A PSF photometry tool for NASA's Kepler, $\mathcal{K}$2, and TESS missions

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Introduction

- NASA's Kepler and K2 missions have been delivering high-precision time series data for a wide range of stellar types
- However, both the official and community developed pipelines [1–3] tend to focus on studying isolated stars using simple aperture photometry, hence performing sub-optimally in crowded fields
- Although Point Spread Function (PSF) photometry methods are well known [4, 5], as of now, there exist no open source tool that takes into account the caveats inherent in Kepler and $\mathcal{K}$2 data
- To address this issue, we present an open source PSF photometry toolkit for Kepler and K2 and extensible to TESS, as part of the $\mathcal{K}$2 Guest Obverser Office data analysis tool, PyKE

Methods

Fitting multiple PSFs jointly

Consider an experiment that outputs an image with m (m known) stellar objects as a collection of n independent *non-identically* distributed random variables $\boldsymbol{Y} \triangleq \{Y_i\}_{i=1}^{n}$ (pixels), each of which has expected value $\mathbb{E}[Y_i] = \sum_{j=1}^{m} \lambda_i(\boldsymbol{\Theta}_j)$, where λ_i is the PSF model at the i-th pixel, $\boldsymbol{\Theta}_j$ is a random vector that encondes the information about flux and center position of the j-th star. Hence, the likelihood function can be written as

$$P\left(\boldsymbol{Y}=\boldsymbol{y}\,\middle|\,\{\boldsymbol{\Theta}_j\}_{j=1}^{m}=\{\boldsymbol{\theta}_j\}_{j=1}^{m}\right)=\exp\left(-\sum_{i=1}^{n}\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)\right)\prod_{i=1}^{n}\frac{\left(\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)\right)^{y_i}}{y_i!}. \quad (1)$$

Perhaps of more practical interest is the log likelihood function

$$\log P\left(\boldsymbol{Y}=\boldsymbol{y}\,\middle|\,\{\boldsymbol{\Theta}_j\}_{j=1}^{m}=\{\boldsymbol{\theta}_j\}_{j=1}^{m}\right)=\sum_{i=1}^{n}\left(-\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)+y_i\log\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)\right). \quad (2)$$

Hence, the Maximum Likelihood Estimator (MLE) can be formulated as the following optimization problem

$$\boldsymbol{\theta}^{*}(\boldsymbol{y})=\underset{\boldsymbol{\theta}\in\Lambda}{\arg\min}\sum_{i=1}^{n}\left(\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)-y_i\log\sum_{j=1}^{m}\lambda_i(\boldsymbol{\theta}_j)\right). \quad (3)$$

Furthermore, prior probability densities (often taken to be uniform on Λ) on the stars positions, fluxes, and sky background, are used.

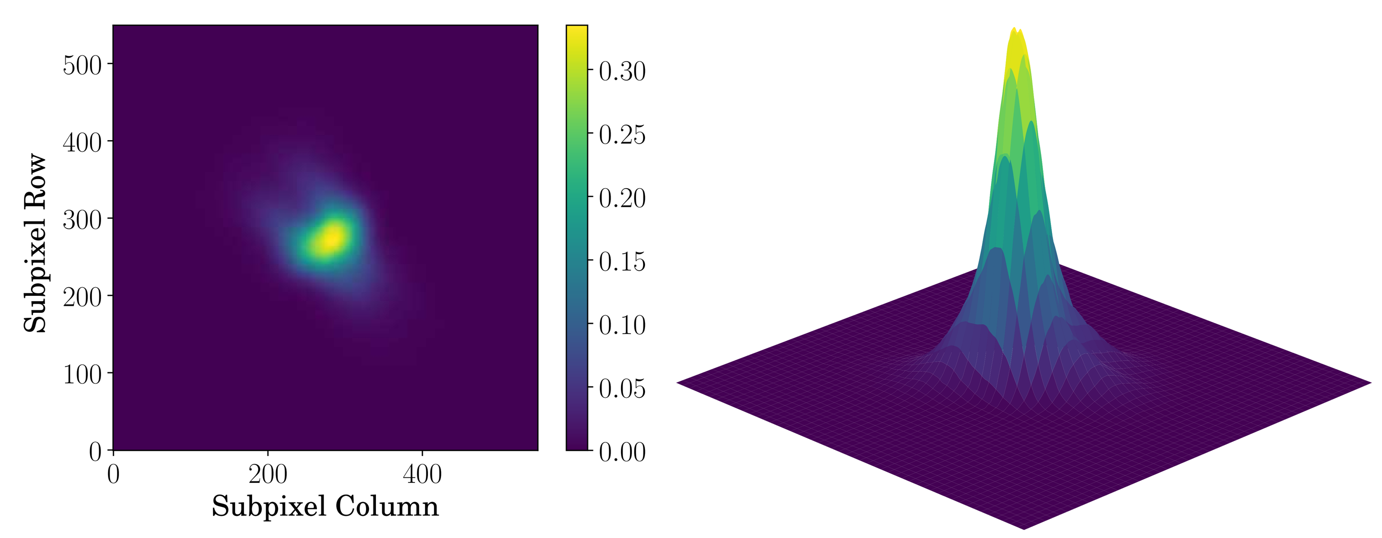
Estimating uncertainties using the Cramér-Rao Lower Bound

Uncertainties on the fitted values are computed using the Cramér-Rao Lower Bound. Mathematically,

$$cov(\boldsymbol{\theta}^{*}(\boldsymbol{Y}))\leq\left(\mathbb{E}_{\boldsymbol{\theta}}\left[\nabla_{\boldsymbol{\theta}}\log p(\boldsymbol{Y}|\boldsymbol{\theta})\left[\nabla_{\boldsymbol{\theta}}\log p(\boldsymbol{Y}|\boldsymbol{\theta})\right]^{T}\right]\right)^{-1}\Big|_{\boldsymbol{\theta}=\boldsymbol{\theta}^{*}(\boldsymbol{y})} \quad (4)$$

The Kepler Pixel Response Function Model

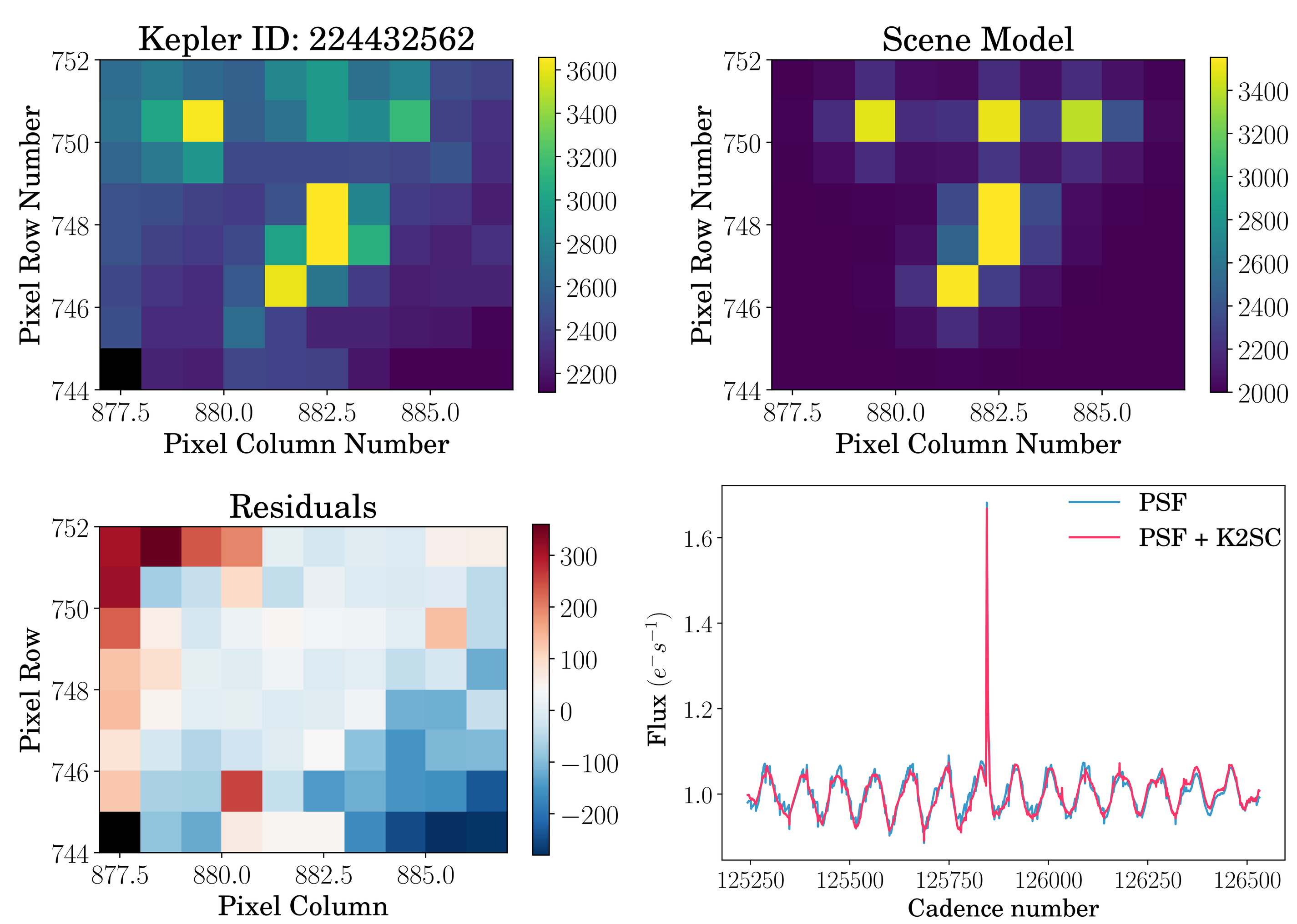
- Kepler's pixel response function (PRF) has been shown to be nonsymmetric and spatially variable accross the detector [6]
- The PRF model used in PyKE is constructed following a similar procedure as stated by Bryson *et al.* [6]



An object of the PSF model can be instantied as follows

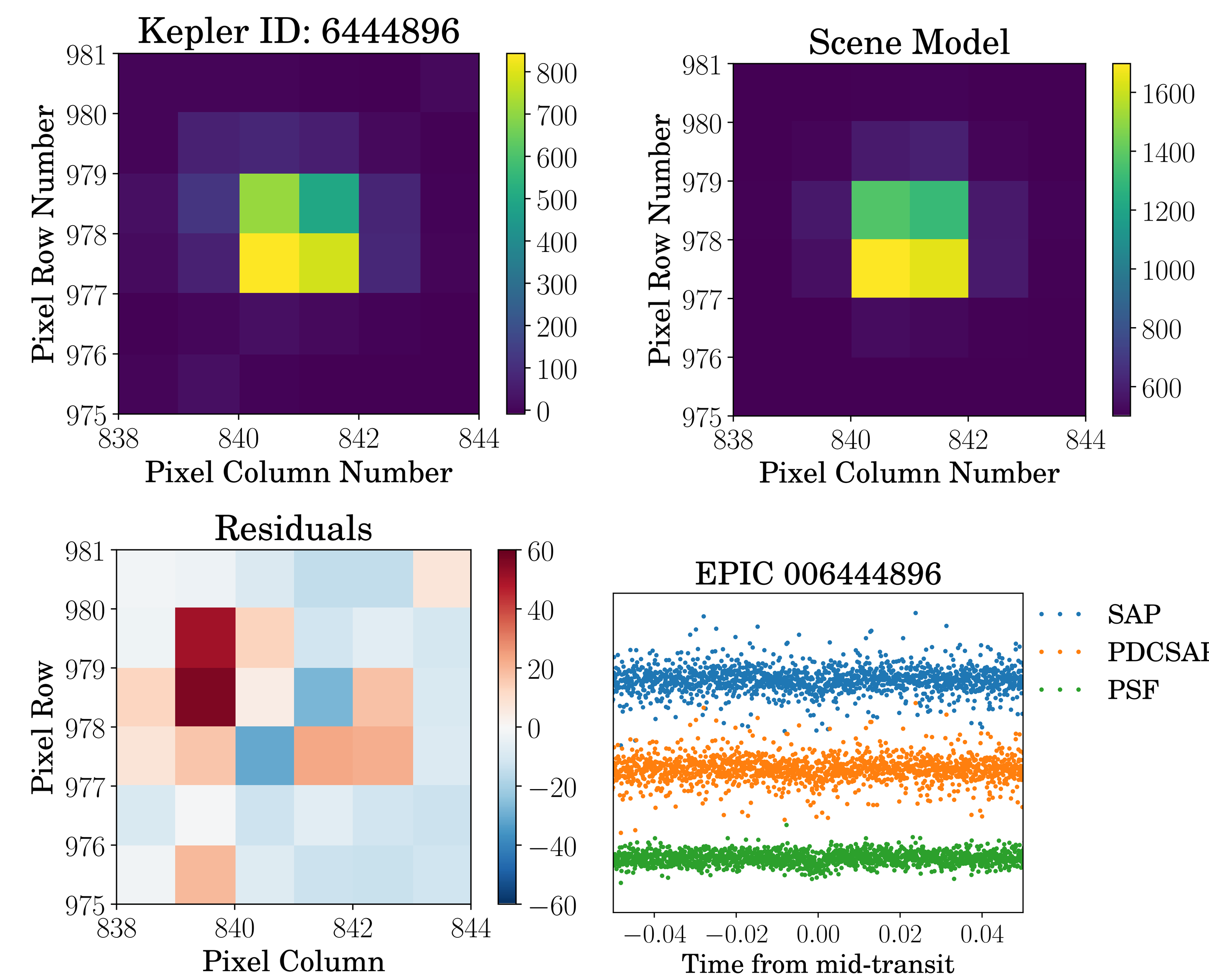
```
>>> tpf = KeplerTargetPixelFile("ADDRESS_TO_TPF")
>>> prf = tpf.get_prf_model()
```

Crowded K2 Clusters




```
>>> scene = SceneModel(prfs=5*[prf])
>>> phot = PRFPhotometry(prior=prior, scene_model=scene)
>>> results = phot.fit(tpf.flux)
```

Kepler Faint Stars



Conclusions

- We have presented an open-source tool to perform PSF photometry on Kepler and $\mathcal{K}$2 data which precisely and accurately estimates stellar positions and fluxes on the CCD
- Motion-dependent noise, primarily caused by the subpixel flat-field variations, which are not taken into account in the PSF model, is still present. Nevertheless, it can be easily removed by the procedures developed by Luger *et al.* [1], Vanderburg *et al.* [2], and Aigrain *et al.* [3]. on the recovered light curves, but it can be easily removed
- As future works, we intend to study novel ways to build the PSF model, especially by using data-driven approaches

References

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