

Fermi-LAT Observations of Extended Gamma-Ray Emission in the Direction of SNR G150.3+4.5

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ABSTRACT

We report here a dedicated analysis of the γ -ray emission around supernova remnant (SNR) G150.3+4.5, observed with the Large Area Telescope (LAT) on board the *Fermi Gamma-Ray Space Telescope*. The Second Catalog of Hard *Fermi* LAT Sources reported detection of a hard spectrum, spatially extended source from 50 GeV - 2TeV, partially overlapping G150.3+4.5. Lowering the energy threshold to 1 GeV we significantly detect a large ($\sigma = 1.40^\circ \pm 0.03^\circ$) extended γ -ray source consistent with the entirety of the radio shell and displaying a power law spectral index of 1.88. [JAM: all these numbers need to be finalized with the gtlike analysis] An obtained HI spectrum toward the SNR suggests that the remnant could be one of the closest to us and estimates of its age indicate that G150.3+4.5 may be in the Sedov-Taylor phase [JAM: not sure about this, maybe distance is very uncertain, and hence age as well? My age estimate was ~ 6 kyr. Need to think about dynamically young vs. young, fast shocks like which SNR? does dynamically young say something more about the progenitor explosion or surrounding medium?]. In contrast, the spectrum of the γ -ray source is more akin to that of a young, leptonic dominated SNR, although ROSAT X-ray observations show no signs of nonthermal emission coincident typically observed in young SNRs. We discuss alternate origin scenarios for the γ -ray emission... [JAM: Should I have the words Pass 8 here somewhere, make it all shorter? move the on board stuff to intro.]

Subject headings: Supernova Remnants, γ -rays, Cosmic rays, Radio

1. Introduction

Something about SNRs, cosmic ray accelerators, radio detections, connection between radio-LAT observations, G150 detection, 2FHL blind detection and SNRs at TeV (all young?), this paper extends the energy down to

We describe the LAT and analysis results in §2, detail multiwavelength observations in §3, and discuss various emission origin scenarios in §4.

2. *Fermi* LAT Observations and Analysis

2.1. Data Set and Reduction

Fermi LAT is a pair conversion telescope sensitive to high energy γ -rays from 20 MeV to greater than 1 TeV (Ackermann et al. 2016), operating primarily in a sky-survey mode which views the entire sky every 3 hours. The LAT has wide field

of view (~ 2.4 sr), a large effective area of ~ 8200 cm² above 1 GeV for on axis events and a 68% containment radius angular resolution of $\sim 0.8^\circ$ at 1 GeV. For further details on the instrument and its performance see Atwood et al. (2009) and Ackermann et al. (2012).

In this analysis, we analyzed 7 years of Pass 8 data, from August 2nd 2008 to August 2nd 2015. The Pass 8 event reconstruction provides a significantly improved angular resolution [JAM: this is sadly unimportant unless I'm at higher energy or using the PSF types. The P8 total PSF at 1 GeV is about the same as for P7REP. It's the acceptance/effective area that are considerably better at this energy], acceptance, and background event rejection (Atwood et al. 2013a,b), all of which lead to an increase in the effective energy range and sensitivity of the LAT. Source class events were analyzed within a $14^\circ \times 14^\circ$ region centered on SNR

G150.3+4.5 using the P8R2_SOURCE_V6 instrument response functions, with a pixel size of 0.1° . To reduce contamination from earth limb γ -rays, only events with a zenith angle less than 100° were included.

For spectral and spatial analysis we utilized both the standard *Fermi* Science Tools (version 10-01-01)¹, and the binned maximum likelihood package *pointlike* (Kerr 2010). *pointlike* provides methods for simultaneously fitting the spectrum, position, and extension of a source, and was extensively validated in Lande et al. (2012). Both packages fit a source model, the Galactic diffuse emission, and an isotropic component (which accounts for the background of misclassified charged particles and the extragalactic diffuse γ -ray background)² to the observations. In this analysis, we used the standard Galactic diffuse ring-hybrid model, scaled for Pass 8 analysis, *gll_iem_v06.fits*, and for the isotropic emission, we used *iso_P8R2_SOURCE_V6_v06.txt*, extrapolated to 2 TeV as in Ackermann et al. (2016).

In our source model for the region, we included sources from the third *Fermi* LAT catalog (Acero et al. 2015, 3FGL) within 15° of the center of our region of interest (RoI). We replaced the position and spectrum of any 3FGL pulsars in the region with their corresponding counterpart from the LAT 2nd pulsar catalog (Abdo et al. 2013). Residual emission unaccounted for by 3FGL sources is present in the RoI due to the increased time range and different energy selection with respect to that in 3FGL. We added to the RoI several point sources to account for this unmodeled emission and minimize the global residuals. [JAM: do I need to say more about these sources? should I mention adding them automatically and iteratively based on TS maps and reference SNRcat/2FHL? How close is the closest source? Mention this and use as an argument for not saying much more about them]. The normalization and spectral index of sources within 5° of the center of the RoI were free to vary, whereas all other source parameters were fixed. A preliminary likelihood fit of the RoI was performed, and sources with a test statistic (TS) < 9 (TS is defined as, $TS = 2 \log(\mathcal{L}_1/\mathcal{L}_0)$ where \mathcal{L}_1

is the likelihood of source plus background and \mathcal{L}_0 that of just the background) were removed from the model.

2.2. Morphological Analysis

Studying the spatial extension of sources with the LAT is non-trivial due to the energy-dependent point spread function (PSF) and strong diffuse emission present in the Galactic plane. To strike a balance between the best angular resolution and minimal diffuse contamination, we restrict our morphological analysis to energies between 1 GeV and 1 TeV. We divide this energy range into 12 [JAM: 4bpd] logarithmically spaced bins for both *pointlike* and *gtlike* binned likelihood analyses.

Three 3FGL sources are located within the extent of G150.3+4.5. 3FGL J0425.8+5600, located approximately 0.6° from the center of the SNR, is the closest of the three sources and is described with a power law spectrum of index $\Gamma = 2.35 \pm 0.17$ in the 3FGL catalog. The closest radio source to 3FGL J0425.8+5600 is NVSS J042719+560823, at 0.25° away (Ref?). 3FGL J0423.5+5442, exhibits a power law spectral index, $\Gamma = 2.63 \pm 0.15$, with no clear multiwavelength source association. Finally, 3FGL J0426.7+5437 has a pulsar-like spectrum, yet it is located about 0.8° from the center of the SNR. In a timing survey performed with the 100-m Effelsberg radio telescope, Barr et al. (2013) were unable to detect pulsations from the source down to a limiting flux density of ~ 0.1 mJy. We discuss this source and potential emission scenarios further in §4.2).

In our analysis, we removed 3FGL J0425.8+5600 and 3FGL J0423.5+544 from the RoI, but kept 3FGL J0426.7+5437 in our model since preliminary analyses showed clear residual emission at the position of the source.

and replaced them with a radially symmetric uniform disk of initial radius $\sigma = 1.5^\circ$

Fill in more about the TS of these sources at 1 GeV prior to removing them and adding in G150 as well as how they're insignificant when added on top of the extended source. Tried fitting and additional extended source that looks like the 2FHL and this was insignificant (I think this negates the need to try splitting the remnant). Different spatial models also not significant at 1 GeV. refit extension

¹<http://fermi.gsfc.nasa.gov/ssc/>

²<http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

of G150 after adding in nearby sources. No templates since the radio we have is too weak.

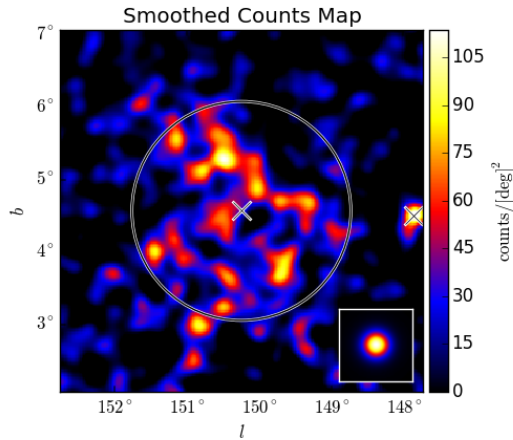


Fig. 1.— Smoothed diffuse and isotropic subtracted counts map, $E > 10$ GeV.

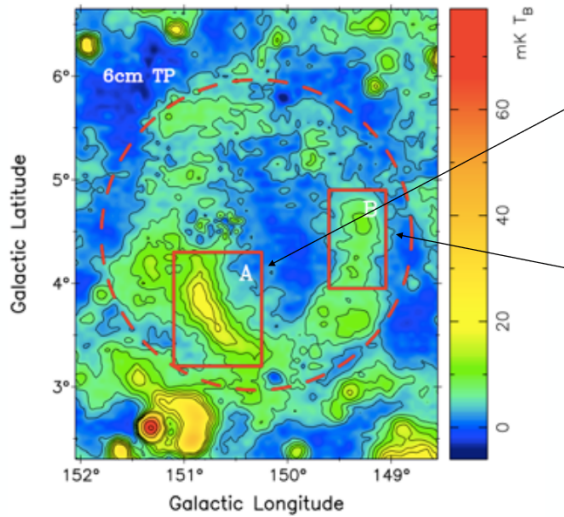


Fig. 2.— Radio image from (Gao & Han 2014)

2.3. Spectral Analysis

Describe gtlike results, spectral models tested (broken PL? no need to since it looks so power law esque?). No break observed, hard spectra increasing to TeV

how many energy bins

2.4. Systematics

Bracketing IRFs and diffuse systematics study still need to be done

3. Multiwavelength Observations and Analysis

Not sure yet if I'll need separate sections

3.1. Radio

I don't think we're presenting any new Radio analysis, just rehashing previous results, showing radio maps overlaid on GeV, so maybe this is really discussion. (Gao & Han 2014)

3.2. HI

3.3. CO

Make CO overlay maps for the possible velocities. Only issue is that Dame 2001 only goes up to 5 deg. Other CO data that covers better to use? Planck?

3.4. X-ray

No diffuse nonthermal X-ray emission observed by ROSAT. No point sources near the center? Should a pulsar be near the center? How to quantify this? Can we place a limit on something like density with an upper limit on X-ray emission? What about other x-ray telescopes?

4. Discussion and Results

4.1. What is it?

Size + HI suggest that near distance corresponding to different HI velocities suggest it's aged, spectrum looks more like young SNR (hard + no GeV break). Is it a weird young remnant or weird aged one? Leptonic dominated if young, hadronic dominated if older? Something about nearby dense clouds masking hadronic emission? Maybe this is only true for MeV cosmic rays that are screened out though and it would only mask the pion bump, but not this higher energy emission?

PWN or SNR. Can we rule out PWN? See W41 paper, MSH 11-61A, Fabios recent G326 work (no,

177 he just tries to use the PSF types and testing dif-
 178 ferent model templates to try to disentangle SNR
 179 from PWN)?

180 No PSR candidate near center (should it be
 181 near the center? Depends on age) Is there some
 182 limit we can place on the PWN based on not
 183 seeing the pulsar? Like on Edot? OR some-
 184 thing like Mattana et al. 2009 correlation between
 185 $\text{flux}_x/\text{flux}_g \propto \text{Edot}$?

186 Assume it's in Sedov phase based on size + near
 187 distance, and calculate age, upper limit on Edot
 188 base on lack of x-ray flux? Or maybe if I assume
 189 the sources is the PWN and GeV radius is PWN
 190 radius, then can I estimate Edot based on size and
 191 evolution inside SNR?

192 If we assume close distance, age is only $\approx 5\text{kyr}$,
 193 maybe this is a transitional SNR? What do others
 194 like this look like? Puppis A? Gamma Cygni is a
 195 similar age too.something

196 4.2. Distance Considerations

197 probably doesn't need to be a different section.

198 4.3. Nonthermal Modeling

199 I think I could get a working model with naima
 200 running pretty quickly, is it worth it?

201 5. Conclusions

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