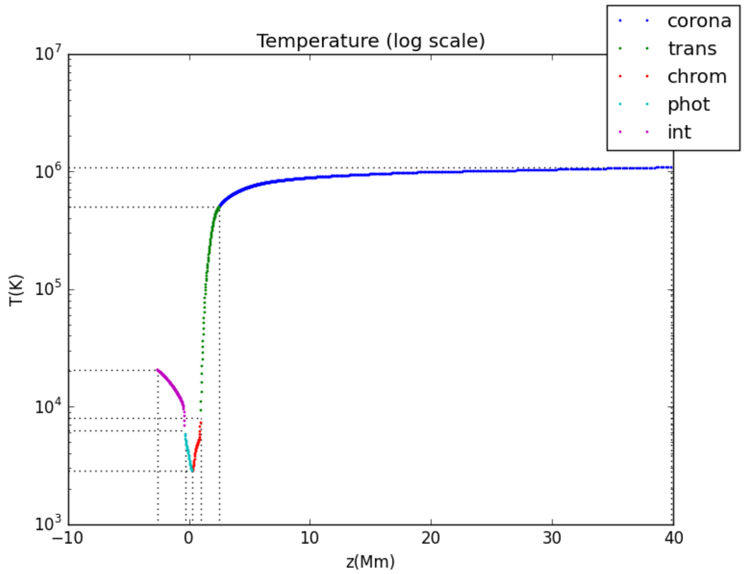

Figura 1: *Temperature vs z plot. logarithmic y scale*

1a) In order to identify the layers I put conditions on temperature:
`http://www.nasa.gov/mission_pages/iris/multimedia/layerzoo.html`

Looking at the values from the file 'atmosphere.dat' ordered by height from top(of the atmosphere) to bottom I consider the corona while temperature ≥ 500000 K (T is decreasing), transition region until T = 8000 K, the chromosphere until T reaches the (only) minimum, afterwards the temperature starts to raise and I consider the layer before it reaches 6500 K the photosphere and the solar interior after

The exact values matching these conditions are:

corona between [39.802200, 2.535930] Mm temperatures: [1.080180e+06, 5.025160e+05] K

transition region between [2.516350, 0.991115] Mm temperatures: [4.991350e+05, 8.067640e+03] K

chromosphere between [0.971556, 0.305708] Mm temperatures: [7.306160e+03, 2.843670e+03] K

photosphere between [0.286093, -0.303487] Mm temperatures: [2.848470e+03, 6.297540e+03] K

solar interior between [-0.323184, -2.592960] Mm temperatures: [6.837750e+03, 2.068340e+04] K

1b) $\mu = \frac{n_H + 4n_{He}}{n_e + n_H + n_{He}}$

$n_H = 10n_{He} \implies \mu = \frac{1.4n_H}{n_e + 1.1n_H}$

- totally ionized H and He $\implies n_e = n_H + 2n_{He} = 1.2n_H \implies \frac{n_H}{n_e} = \frac{5}{6}$ and $\mu = 0.6087$
- neutral H and He $\implies n_e = 0 \implies \mu = 1.2727$

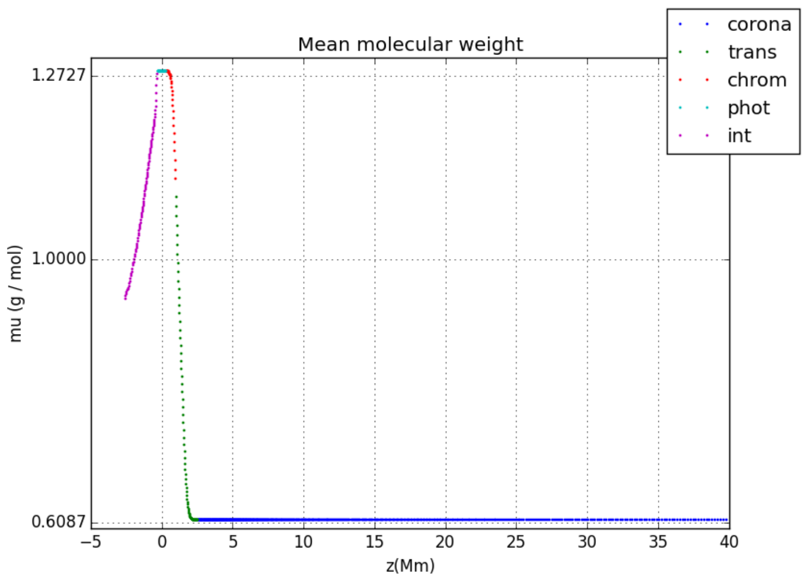

Figura 2: *Mean molecular weight(g/mol) vs z plot* Maximum close to $1.2727 = \mu$ in the case of neutral H and He and minimum close to $0.6087 = \mu$ calculated in the case of completely ionized H and He

$\frac{n_H}{n_e} = \frac{\mu}{1.4-1.1\mu}$

In the case of neutral H and He $n_e \to 0 \implies \frac{n_H}{n_e} \to \infty$

When plotting $\frac{n_H}{n_e}$ using μ from the file, as we can see in the graphic of μ there are some values of z for which $\mu > 1.2727 \implies 1.4 - 1.1\mu < 0 \implies \frac{n_H}{n_e} < 0$

I will limit oy axis values to $[0, 4]$ in order to avoid these negative values and the big ones

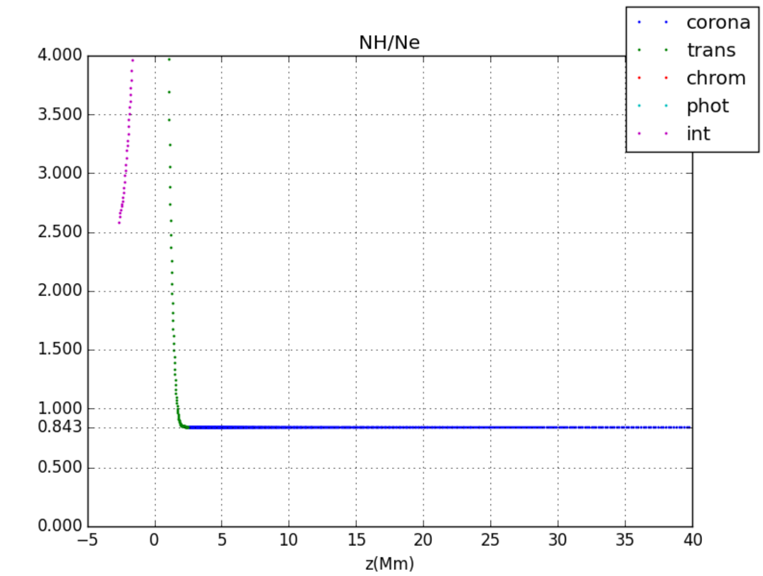

Figura 3: number of atoms of H / number of electrons

We can see a constant value of $\frac{n_H}{n_e}$ in the corona of $\frac{n_H}{n_e} = 0.843 \approx \frac{5}{6}$ which is the value we calculate in the case of totally ionized H and He and we expect this because of the high values of the temperature in the corona

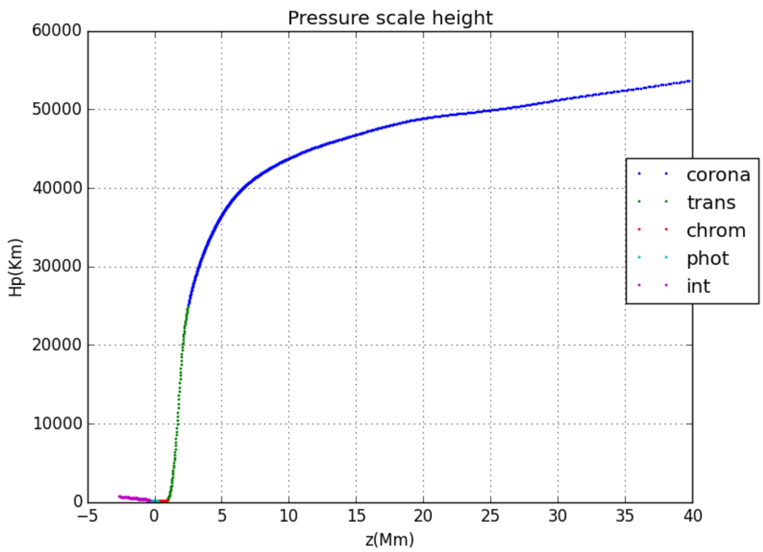

Figura 4: Pressure scale height

$H_p > 0 \implies$ pressure is a decreasing function. H_p is the distance in which pressure will decrease by a factor e so a value close to 0 like in the photosphere and chromosphere means that this distance is very small (it will decrease more abruptly)

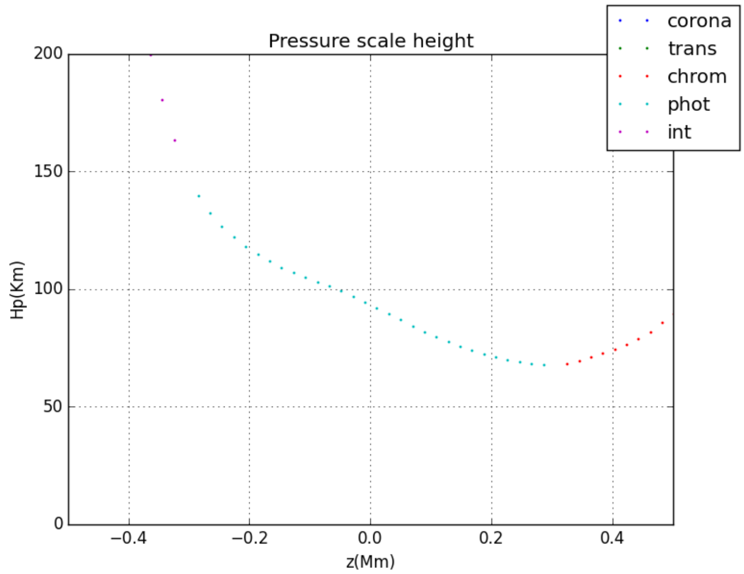

Figura 5: Checking Hp in the photosphere (between approx 90 - 200 km)

2) $\frac{d\ln p}{dz} = -\frac{1}{H_p}$, H_p const $\Longrightarrow$ ln $p(z)$ - ln $p(z_0) = -\frac{1}{H_p}(z - z_0) \Longrightarrow p(z) = p(z_0)exp(-\frac{z-z_0}{H_p})$

$\rho(z) = \frac{1}{gH_p}p(z) = \frac{p(z_0)}{gH_p}exp(-\frac{z-z_0}{H_p}) = \rho(z_0)exp(-\frac{z-z_0}{H_p})$

Analytic test for H_p constant (with values 1 and $1e10$) with $\rho(z_f)$ taking values: $1e-10, 1e-5, 1e-2, 1, 1e2, 1e3, 1e7, 1e10$
Integrating downward or forward in height makes no difference (using ln p)

We see that analytic solution matches exactly numerical solution(we plot ln $p(z)$ - ln $p(z_i)$ vs z) and

that the graphic is a line (expected) with slope:

$\frac{lnp(z_f)-lnp(z_i)}{z_f-z_i} = -\frac{1}{H_p}$

where $z_f = z_{max}$(z at the top of the atmosphere) and $z_i = z_{min}$(z at the bottom of the atmosphere)

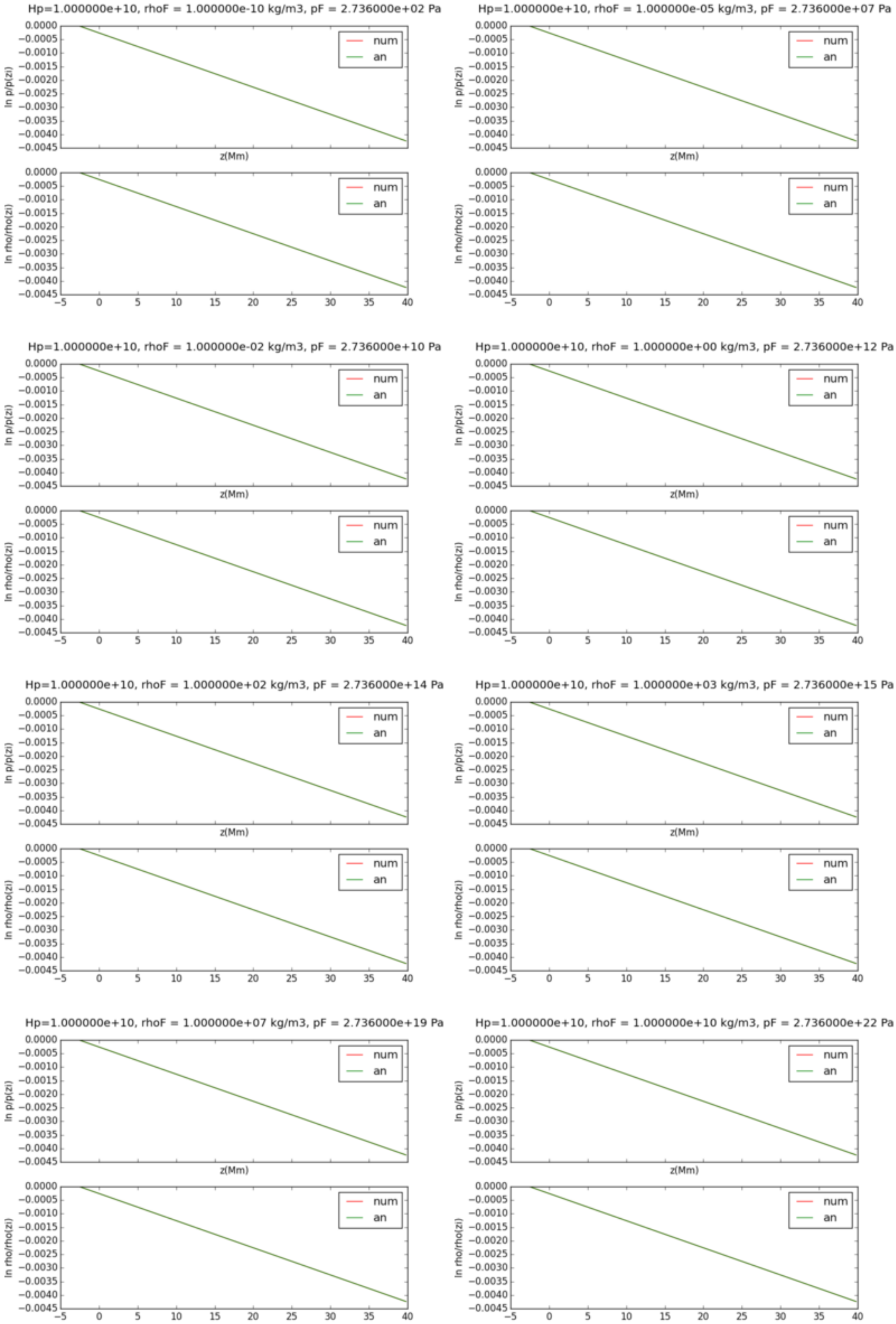

Figura 6: Analytic test Hp=1

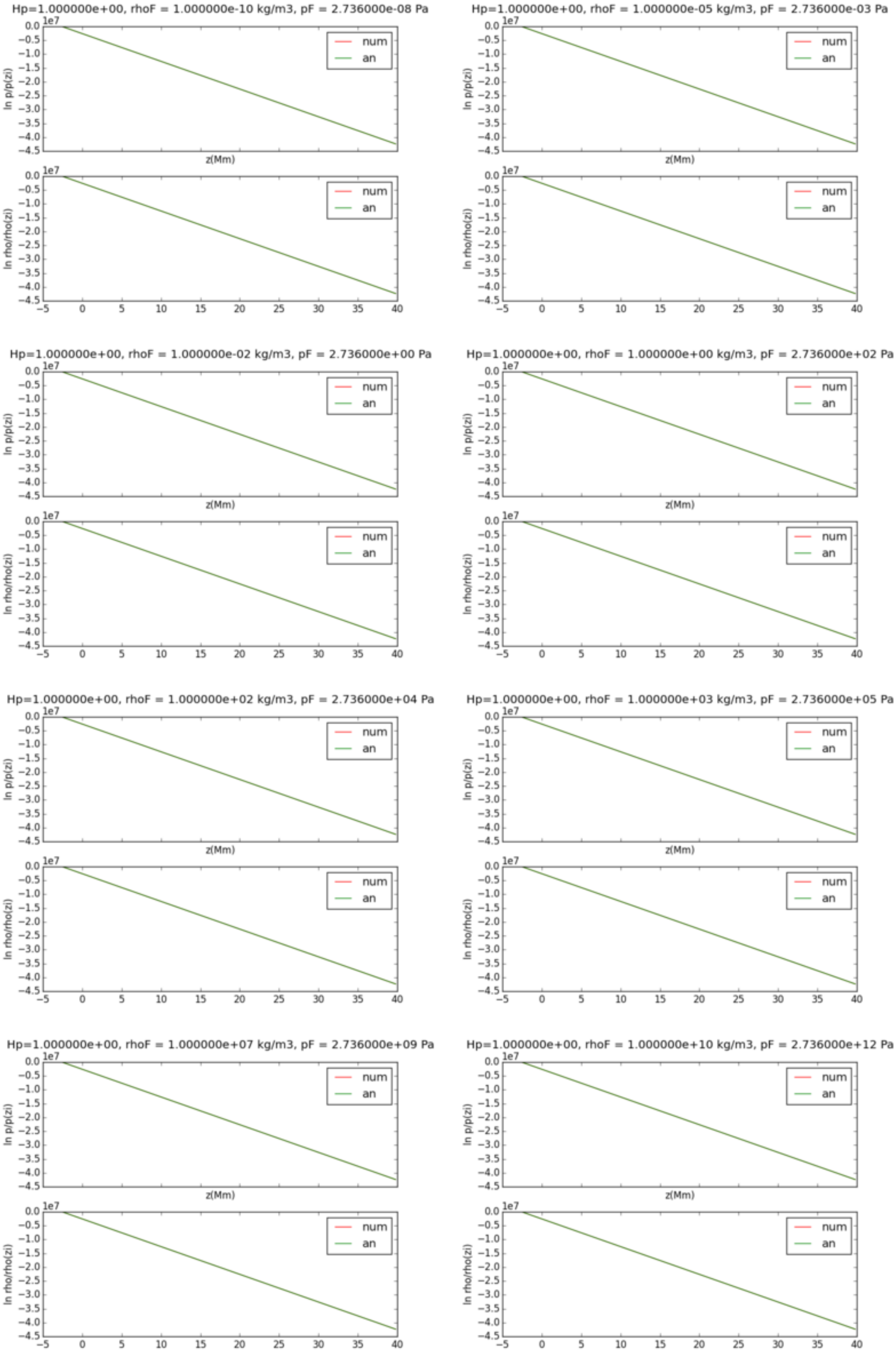

Figura 7: Analytic test Hp=1e10

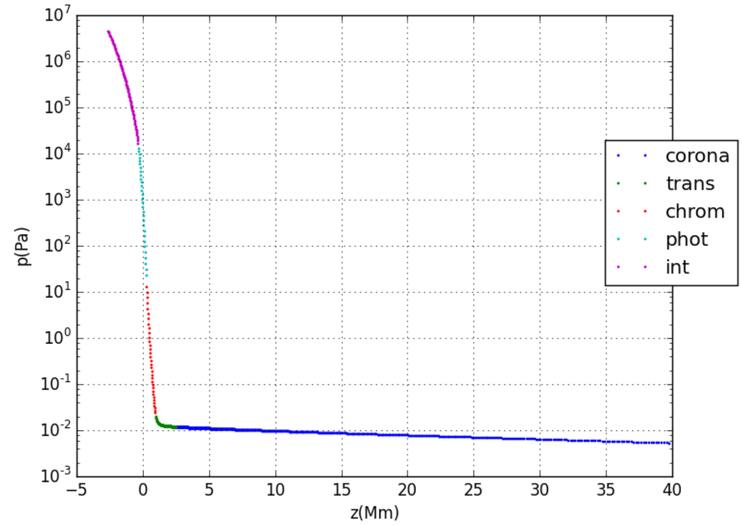

Figura 8: pres log10 oy scale

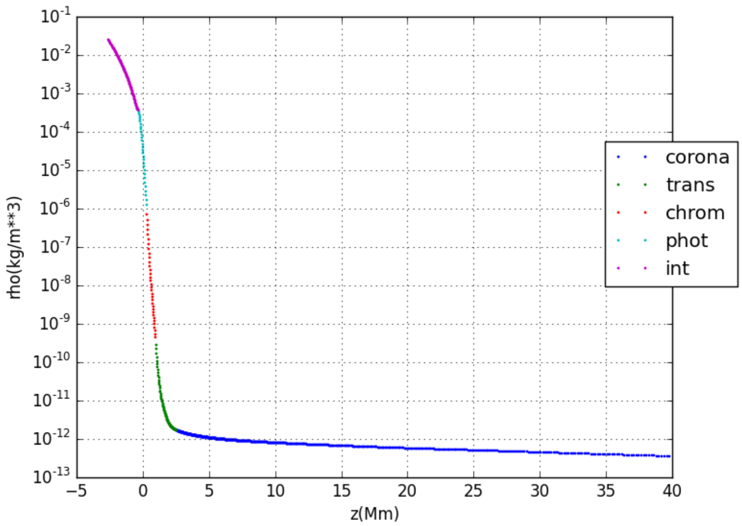

Figura 9: rho log10 oy scale

Pressure will decrease very abruptly in photosphere and chromosphere because H_p is very small in this portion and pressure has still high values

In the transition zone and corona pressure won't decrease fast because H_p has now bigger values and pressure smaller values

But density will decrease fast in the transition zone as well beacuse of the abrupt raise of the temperature and $\frac{p\mu}{\rho T}$ is constant

Notation: μ_0 = magnetic permeability

$\beta = \frac{p}{p_{mag}}$ where $p_{mag} = \frac{B^2}{2\mu_0}$

$v_A^2 = \frac{B^2}{\mu_0 \rho}$

$c_s = \sqrt{\frac{\gamma p}{\rho}}$

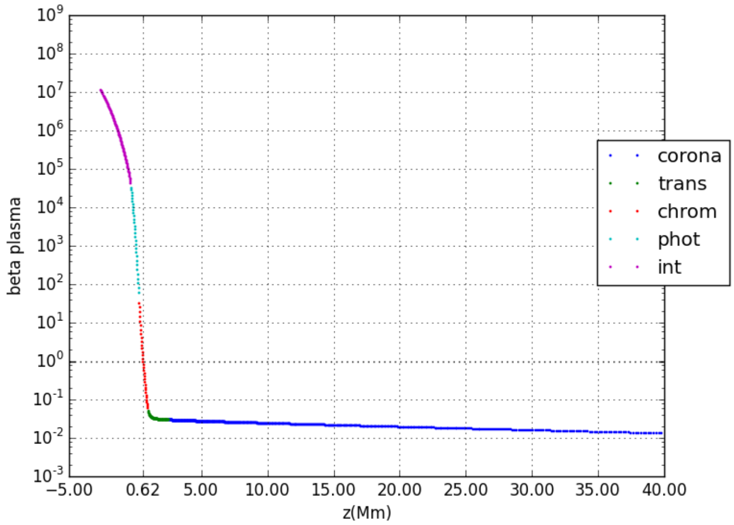

Figura 10: plasma beta log10 oy scale

Plasma beta is a decreasing function and has value 1 at z = 0.62 Mm (in the chromosphere)

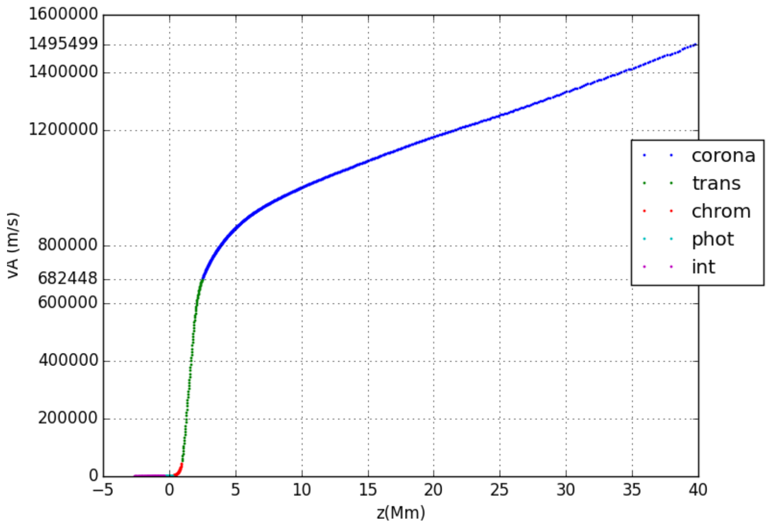

Figura 11: vA

In the corona we observe big values of vA (between approx. 700 - 1500 km /s)

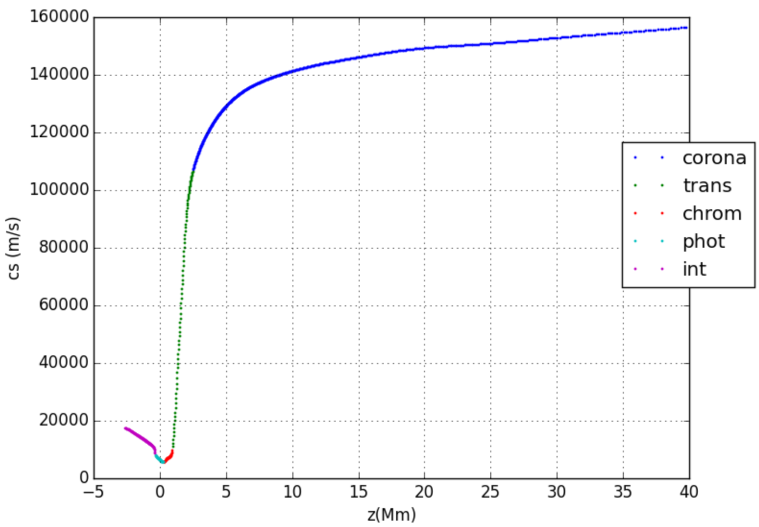

Figura 12: cs

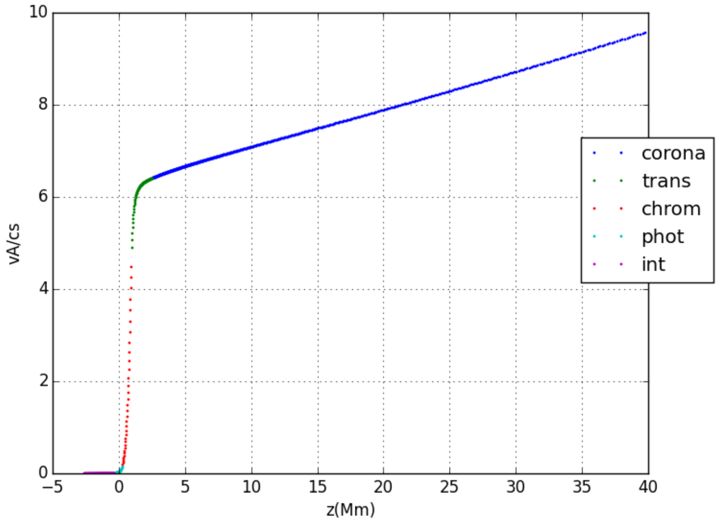

Figura 13: vA/cs

In the corona $v_A > c_s$
$\beta = \frac{2p\mu_0}{B^2} = \frac{2p}{\rho v_A^2} = \frac{2}{\gamma}(\frac{c_s}{v_A})^2 \implies \beta(\frac{v_A}{c_s})^2\frac{\gamma}{2} = 1$ We call this function func(β, $\frac{v_A}{c_s}$) in the graphic below and expect it to be 1

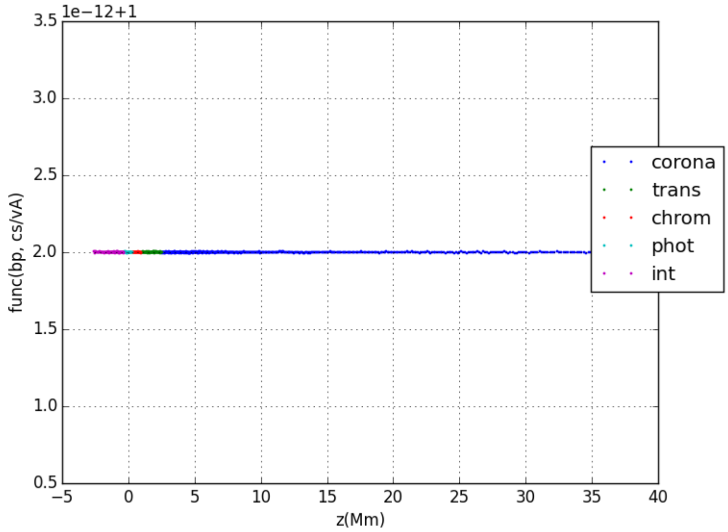

Figura 14: func(bp, vA/cs) $\approx$ 1

3a) units of Λ in c.g.s are $\frac{erg}{cm^3 s}$

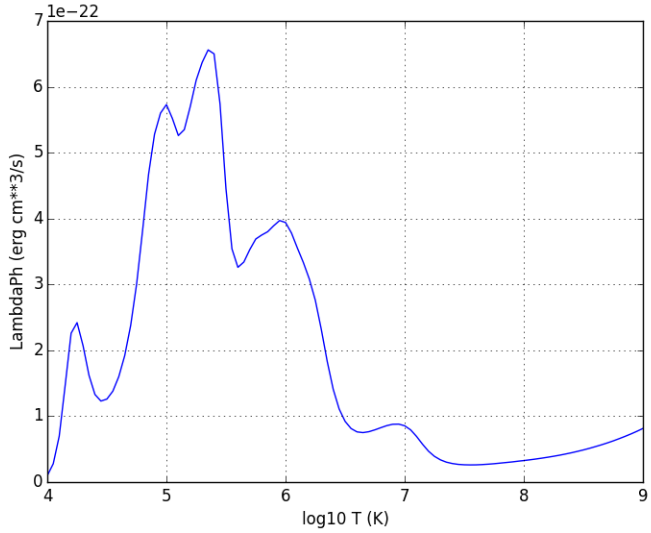

Figura 15: Lambda phot

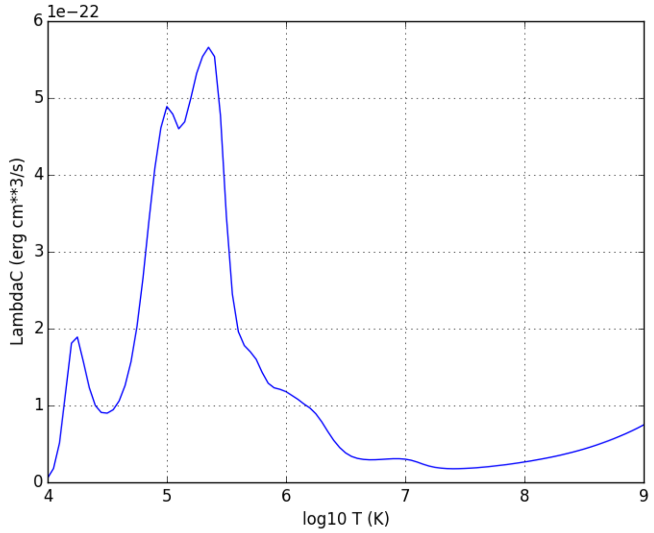

Figura 16: Lambda corona

Both functions have the maximum for T = 2.238721e+05 K

3b) $\rho = \sum_i n_i a_i m_H = (n_H + 4n_{He})m_H$
$n_H = 10n_{He} \implies \rho = 1.4n_H m_H$

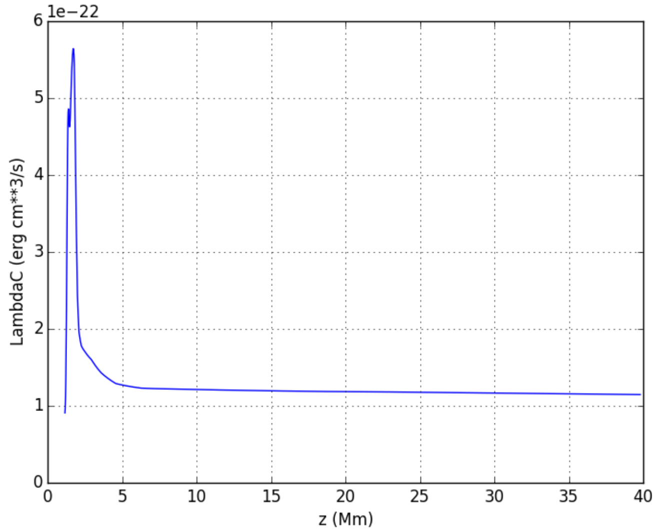

Figura 17: Lambda corona interpolated for atm. temperatures $> 3 * 10^4$ K in 'atmosphere.dat' plotted vs z

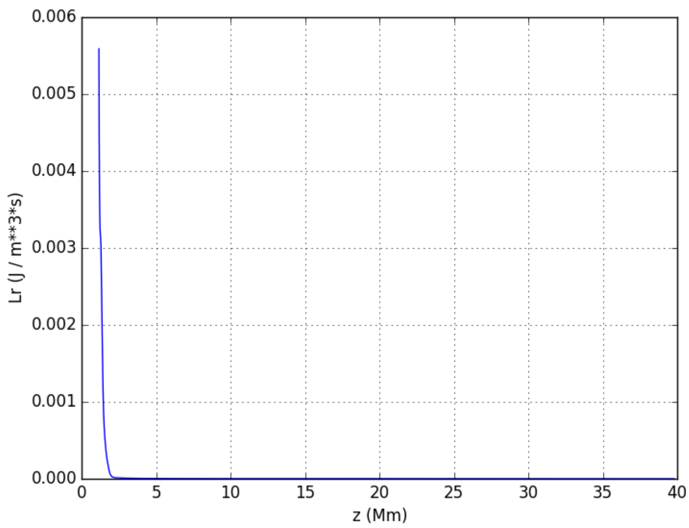

Figura 18: Lr

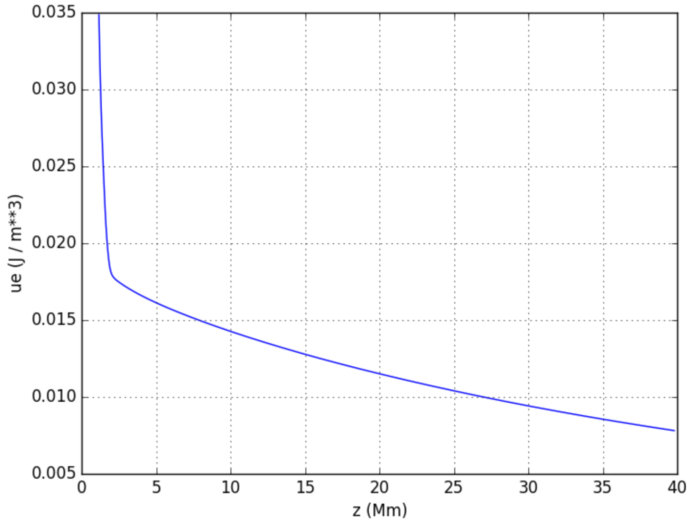

Figura 19: Internal energy

3c)

3d) Equation of energy when $\vec{q} = 0, \vec{v} = 0, \vec{j} = 0$:
$\frac{\partial u_e}{\partial t} = -L_r$
if we consider L_r constant in time:
$u_e(t) = u_e(t = 0) - L_r t \implies \frac{u_e(t=0)}{L_r}$ is the time needed to convert all internal energy into radiation energy
units of $\frac{u_e(t=0)}{L_r}$ are units of time(s)

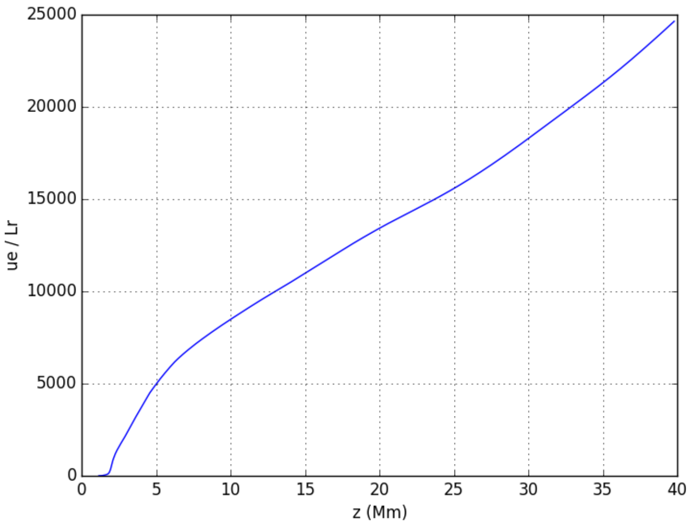

Figura 20: Internal energy / Lr

The maximum value of $\frac{u_e(t=0)}{L_r}$ is about 25000 s in the corona, so in about 7 hours it would lose all internal energy (cool to 0 K ?)