

## Connecting to observations with AGAMA

After including `obs.h` one can use several classes that are useful in connection with observational data. The code resides in six files: `obs_base.h`, `obs_dust.h`, `obs_los.h` and the corresponding `.cpp` files. All objects reside in the `obs` namespace, so in code they should carry the prefix `obs::`.

### *Sky coordinates*

`PosSky` encodes  $(\text{ra}, \text{dec})$  and  $(\ell, b)$  sky coordinates. Possibilities are `PosSky p(40,30)` and `PosSky p(40,30,true)` – the first makes `p` the point with  $\ell = 40, b = 30$  while the second makes `p` the point with  $(\alpha = 40, \delta = 30)$  (degrees). The first component of `p` is `p.l` regardless of whether this actually refers to  $\ell$  or  $\alpha$ , and similarly for the number `p.b`.

`VelSky` encodes a proper motion, so `VelSky v(2,3)` sets the proper motion `v` to  $(\dot{\ell} \cos b = 2, \dot{b} = 3)$  (mas/yr), while `VelSky v(2,3,true)` gives  $(\dot{\alpha} \cos \delta = 2, \dot{\delta} = 3)$ . The components are called `pv.mul` and `pm.mub` regardless of whether they are really  $(\mu_\alpha, \mu_\delta)$  or not.

`PosVelSky` encodes full astrometry – `PosVelSky pv(30,40,2,3)` makes `pv` into a struct that has components `pv.pos`, which is a `PosSky` and thus a component `pv.pos.b`, etc, and `pv.pm`, which is a `VelSky` so has a component `pv.pm.mul`. To store astrometry in the equatorial we write `PosVelSky pv(30,40,2,3,true)` – then `pv.pos.b` will be  $\delta = 40$ , etc. With `p` a `SkyPos` and `v` a `VelSky`, `PosVelSky pv(p,v)` will store in `pv` the astrometry in whichever coordinate system was used for `p` and `v`.

We have a series of translation functions

`PosSky p(from_RAdec(double ra,double dec))` `PosSky p(from_RAdec(PosSky peq))` yield  $(\ell, b)$  positions.

`PosSky peq(to_RAdec(double l,double b))` `PosSky peq(to_RAdec(p))` yield  $(\alpha, \delta)$  positions.

`PosVelSky pv(from_RAdec(ra,dec,mura, mudec))` with `ra,dec,mura,mudec` doubles and `PosVelSky pv(from_RAdec(pveq))` with `pveq`  $(\alpha, \delta)$  astrometry yield astrometry in Galactic coords.

Conversely `PosVelSky pveq(to_RAdec(l,b,mul,mub))` and `PosVelSky pveq(to_RAdec(pv))` with `pv`  $(\ell, b)$  astrometry yield astrometry in equatorial coords.

### *Heliocentric coordinates*

If `intUnits` is an instance of `units::InternalUnits` then `solarShifter sun(intUnits)` creates an object that encodes the position and velocity of the Sun. Invoked thus the Sun's position is  $(x, y, z) = (-8.27, 0, 0.025)$  kpc and its Galactocentric velocity is  $(v_x, v_y, v_z) = (12, 249, 7)$  km s<sup>-1</sup> but these can be changed to anything by creating a `coord::PosVelCar xv` that gives the Cartesian coordinates of the Sun wrt the Galactic centre and then writing `solarShifter sun(intUnits,&xv)`. Observations of the LMC could be handled by making `xv` the position and velocity of the Sun wrt the centre of the LMC.

`coord::PosCar X(sun.xyz())` stores in `X` the Cartesian position of the Sun (internal units). `coord::VelCar V(sun.Vxyz())` stores in `V` the Cartesian velocity of the Sun (internal units).

A `solarShifter` provides several methods for moving between helio- and galacto-centric systems:

If `PosSky p` is in Galactic coordinates and `double sKpc` is a distance in kpc, then `coord::PosCar xyz(sun.toCar(p,sKpc))` stores in `xyz` the Galactocentric Cartesian coordinates associated with `p, sKpc`.

Similarly if `mulb,Vlos` are a proper motion in Galactic coords and a line-of-sight velocity in km/s, we can write `coord::PosVelCar xv(sun.toCar(p,sKpc,mulb,Vlos))`. We can also write `coord::PosVelCyl Rv(sun.toCyl(p,sKpc,mulb,Vlos))` to get the corresponding cylindrical coordinates.

Given a 3d point `p` in either cylindrical or Cartesian coords, the inverse `PosSky lb(sun.toSky(p,sKpc))` returns sky position `lb` (degrees) and distance `sKpc` (kpc). Similarly, given 6d coordinates `pv` (in either cylindrical or Cartesian coords) `PosVelSky lbpm(sun.toSky(pv,sKpc,Vlos))` returns  $(\ell, b)$  astrometry, distance and line-of-sight velocity (kpc, km/s). If `pv` gives 6d coords in either cylindrical or Cartesian coords, we can extract just the proper motion and line-of-sight velocity: `VelSky mulb(sun.toPM(pv,Vlos))`.

If `p` is a 3d position in either Cylindrical or Cartesian coords, double `s=sun.sKpc(p)` stores in `s` the distance to `p` (kpc).

### Dust

A `BaseDustModel` stores a rule for returning the dust density at any location, which probably suffices for observations of external galaxies. `BaseDustModel dsty(5,0.1,norm,12,2,fromKpc)` will create an exponential dust layer with scale length  $R_d = 5$  kpc, scale-height  $z_0 = 0.1$  kpc that has an integral sign ( $m = 1$ ) warped starting at  $R_w = 12$  kpc that reaches to  $\pm h_w = 2$  kpc. The dust density can be scaled up or down by adjusting the double `norm`, while `fromKpc = intUnits.fromKpc` is the linear scaling.

If `potential::PtrDensity gasDens` specifies a gas density, then `BaseDustModel dst(gasDens,norm,fromKpc)` makes the dust density `norm` times the gas density.

If `p` is a 3d position in cylindrical coordinates, double `rho = dsty.dens(p)` store in `rho` the dust density at `p`.

From `BaseDustModel` we derive a more complex class `DustModel` for observations of our Galaxy. The creators of this class have a `solarShifter` as an argument and in addition to one of the global dust patterns inherited from by `BaseDustModel` one can add any number of blobs, spirals or clouds.

`DustModel dsty(Rd,z0,dAvds,Rw,Hw,&sun)` or `DustModel dsty(gasDens,dAvds,&sun)` make `dsty` a `DustModel` that's just the corresponding `BaseDustModel` with a specified location of the Sun and the dust density normalised to `dAvds`  $A_V$  mag per kpc at the Sun.

`dsty.dens(p)` with `p` a `PosCyl` returns the dust density at `p`.

`dsty.addBlob(p,Ampl,rad)` adds a Gaussian blob of dust centred on `p`, `Ampl` and `rad` being the amplitude and scale radius of the blob:

$$\rho(R, z, \phi) = A \exp \left[ \{ (R - p_R)^2 + y^2 + (z - p_z)^2 \} / 2 \text{rad}^2 \right]$$

where  $y \equiv p_R(\phi - p_\phi)$ .

`deleteBlob(n)` deletes the  $n$ th blob – `deleteBlob()` deletes the blob added last.

`addSpiral(Ampl,phase,alpha,Narms,kz)` adds a global `Narms` log-spiral

$$\rho(R, z, \phi) = A \cos [\alpha \ln(R) - N_{\text{arms}}(\phi - \phi_0)] e^{-k_z |z|}$$

`deleteSpiral(n)` deletes the  $n$ th spiral – `deleteSpiral()` deleting the last spiral.

`addCloud(Rc,phic,z0,norm,fname)` adds a cloud centred on  $(R_c, \phi_c)$  with (exponential) scale height  $z_0$  and peak density  $\text{norm}$ . The surface density of the cloud is read from the file `fname` and stored in a linear interpolator. The first line of the file should contain `nx,ny,Q,Xmax,Ymax,amax`, the grid size in the radial and azimuthal directions, Toomre's  $Q$ , the  $x, y$  extents covered by the grid and the peak density. The following numbers should be the density at grid points, the outer loop being over  $y$ . A suitable input file is computed by the method of Lulian & Toomre (1966) – see Binney (2020).

`deleteCloud(n)` deletes the  $n$ th cloud, etc.

### *Lines of sight*

The classes for los in our Galaxy and los to distant galaxies are derived from the class `BaseLos`. A `BaseLos` provides methods to convert  $(\ell, b)$  astrometry and a distance to a location `p`. Similarly, it provides methods to convert a phase-space location into a distance and line-of-sight velocity. Finally, if a dust model is specified it can give the extinction in the  $V, B, R, H, K$  bands at distance  $s$ .

With `Los` an instance of any line of sight, the relevant methods are

`Los.xyz(s)` returns the `PosCar` at distance  $s$  (intUnits)

`Los.Rzphi(s)` returns the `PosCyl` at distance  $s$  (intUnits)

`Los.s(p)` returns the distance (intUnits) to `p` (`PosCar PosCyl`)

`Los.sVlos(xv)` returns a `std::pair<double,double>` comprising the distance and  $V_{\text{los}}$  (both in intUnits) of the phase-space point `xv` (`PosVelCar` or `PosVelCyl`)

`Los.A.V(s)` returns the  $V$  band extinction to  $s$  – `Los.A.H(s)` gives the H-band extinction, etc.

`extLos` is a class of los through distant external galaxies. Since this class is derived from `BaseLos`, instances of `extLos` inherit the methods of `BaseLos`. Distances are measured from the point on the los that's closest to the galaxy's centre, so they are often negative.

`extLos Los(xs,ys,incl,D,from_Kpc)` makes `Los` the line of sight that is `xs` (int units) parallel to the the apparent major axis, `ys` (inr units) parallel to the apparent minor axis of a galaxy that has inclination `incl` (degrees) and is `D` Mpc from us. Created thus, there is no dust so all extinctions  $A_V$  etc vanish.

`extLos Los(xs,ys,incl,D,from_Kpc,&dsty)` where `dsty` is a `BaseDustModel` makes `Los` a los along which there is extinction. This creation call causes  $A_V(s)$  to be tabulated and fitted to a cubic spline.

`SunLos` is a class for lines of sight from the Sun, normally through our Galaxy, but it could be used for lines of sight to nearby galaxies such as the LMC or M31.

`SunLos Los(pos,Sun)` with `pos` a sky position and `Sun` a `solarShifter` makes `Los` a dustless los at `pos`.

`SunLos Los(ps,Sun,&dsty)` with `dsty` a `dustModel` makes `Los` the corresponding dusty los.

### *Distribution functions*

A new keyword `PopFile` for `.ini` files makes it possible to specify the photometric properties of the stellar population described by a distribution function (DF). If the file name given after `PopFile` is valid, AGAMA seeks to read from the file the fraction of stars in some band that are brighter than a series of magnitudes (in ever-fainter order) (xx add colours). When a DF is created (`DF_factory.cpp`), the data are read and stored in a cubic-spline interpolator.

## Interfacing with Observables

The following functions use either a `los` or the ability to specify magnitudes to interface with observations. These functions are specified in `galaxymodel_base.h` so in code their names should carry the prefix `galaxymodel::`:

`computeMomentsLOS(model, &Los, rho, mom1, mom2)` with `model` a `GalaxyModel`, `Los` any `los`, `rho` and `mom1` arrays of doubles and `mom2` an array of `coord::Vel2Cyl` will compute  $\langle \rho \rangle(s)$ ,  $\langle \mathbf{v} \rangle(s)$  and  $\langle v_i v_j \rangle(s)$  along `Los`. If `Los` is a `SunLos`,  $\langle \rho \rangle$  includes an  $s^2$  factor, so it has units of mass per sterad per unit distance. If `Los` is an `extLos` there is no  $s^2$  factor and the units are mass per unit volume. The inclusion or not of the  $s^2$  factor affects the amount by which foreground and background stars contribute to the velocity moments but not their units.

`computeMomentsLOS` has several optional parameters: the first three are uncertainties on the three moments, the next, `bool separate`, specifies whether the contributions from the components of `model` (bulge, stellar halo, thin disc, etc) should be reported separately. The next two specify the required relative error and the maximum number of DF evaluations, and the final two, `bright` and `faint`, limit the stars to be considered to a range of *apparent* magnitudes. Apparent magnitudes are connected to absolute magnitudes using the distance and extinction provided by `Los`. For example

`computeMomentsLOS(model, los, &dens[0], &Vphi[0], &sigmas[0], NULL, NULL, NULL, true, 1e-3, 1e5, 8, 15);` with `std::vector<double> dens(nc), Vphi(nc)` and `std::vector<coord::Vel2Cyl> sigmas(nc)` will compute moments, separated into the `nc` components, accurate to 0.1% for stars between mags 8 and 15.

`sampleVelocity(model, Rzphi, N, &Los, 8, 15)` with `Rzphi` a `PosCyl` will return (as `std::vector<coord::VelCyl>`) the velocities of `N` mock stars at the `coord::PosCyl` `Rzphi` that have apparent magnitudes between 8 and 15 given the extinction values provided by `Los`. The contributions of individual components can be separated by calling this function multiple times with `model` defined to include only one component's DF. Since disc DFs vanish for  $v_\phi < 0$ , there is an analogous function `sampleHalfVelocity` that samples the interesting half of velocity space.

`sampleVelocity_LF(model, Rzphi, N, &Los, 8, 15)` returns an apparent magnitude in addition to a velocity for each mock star (as `std::pair<coord::VelCyl, double>`).

`sampleLOS(model, &Los, N, 8, 15)` returns `N` phase-space locations (`std::vector<coord::PosVelCyl>`) for stars between apparent mags 8 and 15. The  $s^2$  factor is included when `Los` is a `SunLos`. The last two parameters are optional; omitting them will remove restriction on apparent magnitudes.

`sampleLOS_LF(model, &Los, N, 8, 15)` returns an apparent magnitude in addition to a phase-space location for each mock star (as `std::pair<coord::PosVelCyl, double>`).

`sampleLOSsVlos(model, &Los, N, 8, 15)` returns just distances and  $V_{\text{los}}$  values for mock stars (as `std::vector<std::pair<double, double> >`).