

## PLANCK CLUSTER PAPER

SB<sup>1</sup>, JPH<sup>1</sup>, FM<sup>2</sup>, PD<sup>1</sup>

*Draft Version July 27, 2018*

### ABSTRACT

**Copied and pasted from proposals.** We propose to continue our program of optical imaging to unveil all of the most massive clusters in the observable Universe. We start from the all-sky *Planck* Sunyaev-Zeldovich (SZ) catalogs, which contain several hundred high significance (signal-to-noise ratio, SNR > 5) unconfirmed cluster candidates. Since SZ selection favors high mass clusters and the *Planck* confirmation process favored low redshift systems, the highest significance unconfirmed candidates are, therefore, likely massive clusters ( $M_{500} > 5 \times 10^{14} M_\odot$ ) at relatively high redshift ( $z > 0.5$ ). Our proposed observations, using MOSAIC-3 on Mayall, are designed to confirm the presence of a brightest cluster galaxy (to  $z \sim 1$ ) and red sequence of accompanying cluster members (to  $z \sim 0.7$ ). Preliminary results from our observations over the past two years have validated our approach by the detection of optical clusters in a number of *Planck* candidates, including the discovery of rich systems at  $z = 0.553$  and  $z = 0.830$  that rival the most massive clusters known. The proposed observations represent the first step required to provide a complete all-sky census throughout the observable Universe of the most massive, high redshift clusters. Their expected high redshift and high mass make the unconfirmed *Planck* clusters, arguably, the most important available sample for probing deviations from  $\Lambda$ CDM and defining the high-mass end of the cluster mass function.

### 1. INTRODUCTION

this section has not been edited and is just a bunch of stuff copy and pasted. I did update some of the references.

Massive clusters of galaxies are the extraordinary objects which hold important clues to the evolution of structure in the universe. The widely accepted  $\Lambda$ CDM model of cosmology makes specific predictions about the mass distribution of galaxy clusters in the universe. The number of galaxy clusters, especially at high redshifts, can help constrain structure formation models (e.g., Mortonson et al. 2011; Harrison & Coles 2012; Harrison & Hotchkiss 2012; Waizmann et al. 2012; Zitrin et al. 2009). Galaxy clusters also harbor a significant fraction of the visible baryons in the Universe, in the form of a hot intracluster medium that leaves an imprint on the Cosmic Microwave Background (CMB) through the Sunyaev-Zel'dovich effect (SZ; Sunyaev & Zeldovich 1972 effect).

Using the SZ effect to discover clusters of galaxies has the distinct advantage that the surface brightness of the SZ effect does not dim with increasing redshift. This allows homogeneous samples of massive clusters to be detected out to arbitrary distances. Ground based, large area-sky surveys such as those with the Atacama Cosmology Telescope (ACT; Swetz et al. 2011) and the South Pole Telescope (SPT; Carlstrom et al. 2011) have produced catalogs of hundreds of massive clusters below  $z \sim 1.4$  (e.g., Hasselfield et al. 2013; ?). Now, *Planck* (Tauber et al. 2010; Planck Collaboration et al. 2011) has released an all-sky SZ sample (PSZ; Planck Collaboration et al. 2014b, 2015b) that contains 861 confirmed clusters (of which most [683] were known previously) and another 366 unconfirmed cluster candidates.

Clusters were initially confirmed by cross correlating with previous catalogs (see Section 4; Planck Collabo-

ration et al. 2014b). More recently, dedicated follow up of still-unconfirmed clusters has begun in earnest (e.g., Liu et al. 2015; Planck Collaboration et al. 2015b, 2016; Burenin 2017; Barrena et al. 2018; Amodeo et al. 2018; Streblanska et al. 2018).

This paper is organized as follows: sections 2 through 4 describe the design, observations, data reduction and calibration, and creation of derived data products. In Section 5, we present the main results of our observations, and discuss the results in Section 6. In Section 7, we summarize the key results and conclude.

Unless otherwise noted, throughout this paper, we use a concordance cosmological model ( $\Omega_\Lambda = 0.7$ ,  $\Omega_m = 0.3$ , and  $H_0 = 70 \text{ km s}^{-1}\text{Mpc}^{-1}$ ), assume a Chabrier initial mass function (Chabrier 2003), and use AB magnitudes (Oke 1974).

### 2. DESIGN

Our observational design is motivated by the release of the second, all-sky PSZ catalog<sup>3</sup> (hereafter PSZ2; Planck Collaboration et al. 2015b) which contains 559 unconfirmed SZ detections with  $S/N > 4.5$ . We posit that the vast majority of these must lie at  $z > 0.4$  because the *Planck* confirmation process (Planck Collaboration et al. 2014b) mostly relied on existing catalogs which have a preference for low- $z$  clusters. Furthermore, the confirmed sample of PSZ2 has only a small fraction (3%) of  $z > 0.6$  clusters compared to that expected ( $\sim 20\%$ ) based on the theoretical halo mass function (e.g., Jenkins et al. 2001; Tinker et al. 2008) for mass limit of XXXX. If other clusters as massive like el “El Gordo” exist, they are hiding as high-significance candidates within the objects in this *all-sky* catalog.

The core of our observational design relies on the use of optical imaging to confirm the SZ detections as real clus-

<sup>1</sup> Rutgers;boada@physics.rutgers.edu

<sup>2</sup> Illinois

<sup>3</sup> [http://szcluster-db.ias.u-psud.fr/sitools/client-user/SZCLUSTER\\_DATABASE/project-index.html](http://szcluster-db.ias.u-psud.fr/sitools/client-user/SZCLUSTER_DATABASE/project-index.html)

ters and provide photometric redshifts using the multi-color information. This design is based on the previous success with the ACT cluster confirmation process using 4-m class telescopes. For example, assuming WMAP7 cosmology (Komatsu et al. 2011) with the Tinker et al. (2008) halo mass function, there should be only  $\sim 4$  clusters as massive as El Gordo ( $\leq 2 \times 10^{15} M_{\odot}$ ) at  $z > 0.6$  in the full area covered by the *Planck* PSZ catalog (83.7% of the sky). Although *Planck*s larger beam size (compared to both ACT and SPT) makes it more sensitive to clusters at lower redshifts (due to their larger projected area on the sky), among the confirmed clusters in the recently released all-sky *Planck* SZ catalog are the two highest significance high-redshift SZ detections from ACT (as well as several other ACT and SPT clusters).

Our strategy for this project is to use the Kitt Peak National Observatory (KPNO) Mayall-4m telescope imaging as the first and fundamental step to confirm the highest significance detections in the PSZ2 catalog that are visible across the entire northern sky. Following closely the procedure used for ACT follow-up (e.g., Menanteau et al. 2013), targets are prioritized by SZ signal-to-noise ratio (SNR). We choose to initially report on targets with PSZ2 SNR  $> 5$  as the statistical reliability of PSZ2 cluster candidates should be quite high: according to the *Planck* team  $\sim 90\%$  of candidates at SNR  $> 5$  are expected be “real” clusters (see Figure 11; Planck Collaboration et al. 2015a). Optical imaging should be sufficient to confirm nearly all of the candidates, but for the highest redshift ones, near-IR data will be necessary. Again following the procedure for ACT cluster follow-up: those candidates with some evidence for a high- $z$  brightest cluster galaxy (BCG) will be targeted with near-IR observations to confirm the presence of a BCG and detect the red sequence of cluster members. Observational priority again is given to higher S/N candidates.

### 2.1. Observations

All observations were conducted with the KPNO Mayall telescope. The optical observations were made with the MOSAIC camera mounted at the prime focus. Two detector packages were used for the observations. The earlier MOSAIC1.1 instrument consisted of eight  $2048 \times 4096$  SiTe CCDs, arranged  $2 \times 4$ , separated by a  $\sim 50$  pixels gap with a pixel scale of  $0''.26$  pixel $^{-1}$ . MOSAIC1.1 was replaced with Mosaic3, in mid-2015, and consists of four new  $4k \times 4k$ , 15 micron pixel, 500-micron thick LBNL deep-depletion CCDs. Because the only change from MOSAIC1.1 to MOSAIC3 are the CCDs and controllers both versions have a  $36' \times 36'$  field-of-view.

The optical observing strategy consists of targeted *griz* observations of individual candidates with total exposure times of 360 s, 360 s, 1100 s and 1100 s (assuming dark conditions). The final exposures consist of four dithered positions with individual exposures of 90 s for the *gr*-bands or 275 s for the *iz*-bands. These exposure times are designed to provide  $5\sigma$  detections limits of  $g = ??$ ,  $r = 24.5$ ,  $i = 24.5$ ,  $z = 24.2$  ensuring the unambiguous detection of the faint (i.e.,  $0.4L\star$ ) galaxies in the red cluster sequence up to  $z \sim 1.0$  (citation?) and of brightest cluster galaxies (BCGs) to higher redshifts. The choice of filters in our program is driven by the need to segregate early-type galaxies in the cluster through their colors (or photometric redshifts) by sampling blue-ward and red-

ward of the  $4000\text{\AA}$  break. Our depths are quite a bit different than the designed depths. Should we mention that here, or wait till later on when we are discussing how we actually did?

## 3. DATA REDUCTION AND CALIBRATION

Standard image reductions including subtraction of dark frames, flat fielding, sky-subtraction, and bad pixel masking was performed by the NOAO virtual observatory using the MOSAIC (Valdes & Swaters 2007) science pipelines. The resultant FITS files consist of fully reduced images with either all single exposure CCDs mosaicked into a single image extension (as in the case of Mosaic1.1) or as a multi-extension FITS file with each single exposure CCD occupying a separate extension.

We then mosaic each separate exposure into a master mosaic as described in the following section.

### 3.1. Mosaicking

Combined mosaics are created with SWARP (Bertin et al. 2002). We create three distinct types of mosaics. The individual dither frames are stacked and then median combined to produce the final completed science mosaic. A “detection” is created by combining select science mosaics into a “chi2” image using either the *i*- and *z*-band when both are available and of sufficient quality. Finally we create a set of mosaics use to produce the three color image used for cluster finding. We median combine the *griz* science mosaics into a “blue” (*g*-band), “green” (*r*-band), and “red” (*iz*-band) mosaic. All final mosaics have a pixelscale of  $0''.25/\text{pix}$ . The final exposure time is calculated as the median exposure time of the combined images, and similarly the final airmass is median of the individual air masses. need to talk about the weight images

The full parameter file used while creating the mosaics is given in Appendix ??.

### 3.2. Source Extraction and Photometry

For source extraction and photometry estimation we use Source Extractor (hereafter SExtractor; version 2.19.5; Bertin & Arnouts 1996) run dual image mode with the CHI2 detection image as the detection image. See Section 3.1. See Appendix ?? for a complete parameter listing.

### 3.3. Astrometric Calibration

Each of the final science mosaics produced in the previous section are first astrometrically aligned with *Gaia* (Gaia Collaboration et al. 2016a) Data Release 1 (Gaia Collaboration et al. 2016b) using SCAMP (Bertin 2006) as a part of PHOTOMETRYPIPELINE<sup>4</sup> (PP; Mommert 2017).

Sources are extracted from the mosaics with a signal-to-noise ratio (SNR) of at least ten and with a minimum area of at least 12 pixels. The extracted sources are then matched to the *Gaia* data and a new astrometric solution is calculated. Because the initial astrometric solution from the VO is quite accurate, the resultant corrections are much less than  $1''$ .

<sup>4</sup> <https://github.com/mommert/photometrypipeline>

### 3.4. Photometric Calibration

After the mosaics have been astrometrically aligned, we use PP to produce a photometric solution. PP calculates a photometric zero-point in each of our observed bands by comparing field stars located throughout the mosaic to known photometry from large-area sky surveys. Because our sources are spread across the entire northern sky, and because we prefer to minimize the number of differences between photometric solutions we are limited to two optical surveys. For the optical data, we first seek photometric data from the *Sloan Digital Sky Survey* (SDSS; York et al. 2000) Data Release 13 (DR13; Albareti et al. 2017). When our target does not lie within the SDSS footprint we utilize the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS; Chambers et al. 2016) Data Release 1 (hereafter PS1; Flewelling et al. 2016). Both surveys provide accurate *griz* magnitudes and large on-line queryable databases for rapid automated calibration.

Sources are extracted from the combined mosaics with either a 3'' or 8'' diameter aperture for optical sources respectively; sources with a SNR  $\geq 10$  are matched to a survey catalog and a photometric zero-point is determined. We use half of the available stars (with accurate catalog photometry) to derive the zero-point resulting in zero-points calculated from approximately 10 – 500 stars and with typical uncertainties of 0.05 mag for the *griz*-bands and 0.16 mag for the *z*-band.

**Should we talk about the difference between us and SDSS? If so, how should we “sum up” the differences in a simple way?**

## 4. ANALYSIS

### 4.1. Photometric Redshifts

We determine photometric redshifts ( $photo-z$ ) from the five-band optical images using Bayesian Photometric Redshifts (BPZ; Benítez 2000; Coe et al. 2006) following the same procedure as in Menanteau et al. (2009).

We assess the effectiveness of our  $photo-z$  estimates by comparing with the available spectroscopic redshifts ( $spec-z$ ) from the SDSS. We use three diagnostics to gauge  $photo-z$  accuracy. First, we report the full scatter between the  $photo-z$  and  $spec-z$ , defined as:

$$\sigma_f = \text{RMS}[\delta z / (1 + z_{spec})] \quad (1)$$

where  $\delta z = z_{spec} - z_{phot}$ . Second, we report the normalized median absolute deviation (NMAD; Ilbert et al. 2009; Dahlen et al. 2013; Molino et al. 2017), given as

$$\sigma_{NMAD} = 1.48 \times \text{median}\left(\frac{|\delta z|}{1 + z_{spec}}\right). \quad (2)$$

which provides an estimate of the scatter resistant to catastrophic outliers. Finally, the catastrophic outlier fraction (OLF) where we define a catastrophic outlier (following Molino et al. 2017) as,

$$\eta = \frac{|\delta z|}{(1 + z_{spec})} > 5 \times \sigma_{NMAD}. \quad (3)$$

Figure 1 shows the  $photo-z$  performance as a function of the true spectroscopic redshift. Because we are primarily concerned with identifying clusters containing elliptical galaxies, we show only galaxies classified E/S0

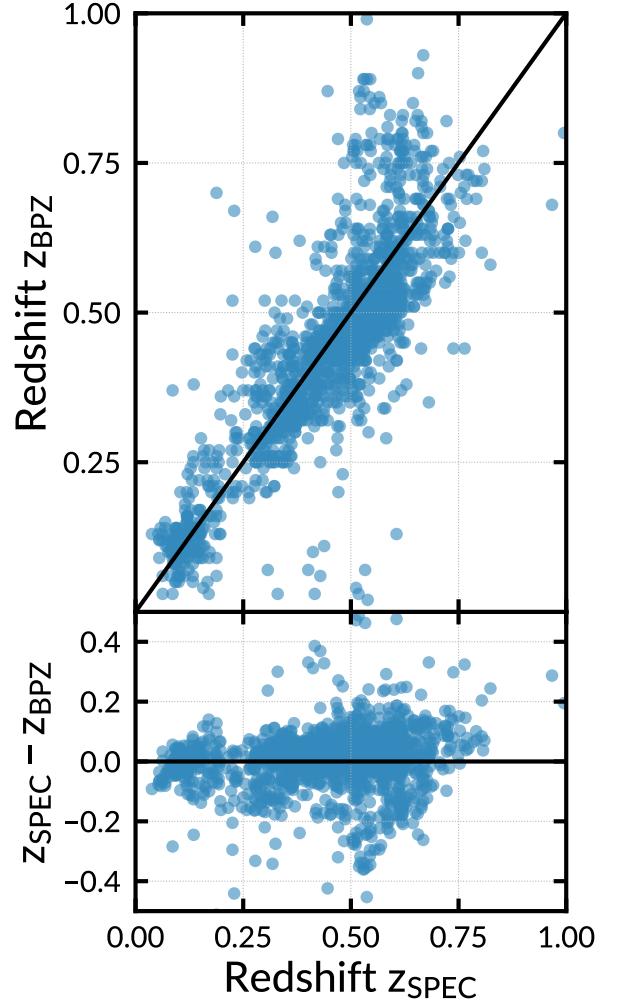


FIG. 1.— Comparison between photometric and spectroscopic redshifts for 1588 elliptical galaxies which have spectroscopic redshifts from the SDSS. The photometric redshifts in the top panel use a Bayesian approach with a custom empirical prior on galaxy brightness for the photometric redshifts. The bottom panel shows the difference between the spectroscopic and photometric redshift.

by BPZ. We find  $\sigma_f = 0.067$ ,  $\sigma_{NMAD} = 0.048$ , and an outlier fraction,  $\eta = 0.9\%$ .

### 4.2. Cluster Finding

In this section, we briefly describe the algorithms and methods used to select the galaxy clusters from the multi-wavelength optical imaging. We follow the methods described in detail in Menanteau et al. (2009, 2010a). We direct the reader there for an in depth description and discussion of the methods.

We first create a three-color image using STIFF (Bertin & Emmanuel 2011). The red, green, and blue channels are given by the corresponding combined mosaics described in Section 3.1. We then visually inspect an area of roughly  $8' \times 8'$  centered on the position of each unconfirmed cluster. Potential brightest cluster galaxies (BCGs) are identified by first calculating the absolute limiting magnitude [needs details](#).

Once a potential BCG is selected, the algorithm selects nearby galaxies, within  $|z_{BCG} - z| < 0.05$  and 0.5 Mpc

projected radius, which BPZ has classified as either E or E/S0 galaxies. These photo- $z$ 's of the galaxies are combined using a  $3\sigma$  median sigma-clipping algorithm to estimate the cluster's mean redshift,  $z_c$ . We use this mean cluster redshift measurement and the member selection criteria given previously to estimate the number of cluster members within 1 Mpc,  $N_{1\text{Mpc}}$ , which we define as the richness of the cluster,  $N_{gal}$ .

We correct the  $N_{gal}$  estimate by subtracting a statistical background of galaxies. We first estimate the number of background ellipticals by selecting galaxies within an annulus ( $R_{200} < r < 2R_{200}$ ) around each cluster's position. We include galaxies with  $\delta z = 0.05$  and similar colors as those galaxies assumed to belong to the cluster. These galaxies are subtracted from the cluster's population which provides an corrected  $N_{gal}$ ,  $N_{galc}$ , which we then use to compute other important quantities. In practice the corrected number of galaxies is between 15% and 20% lower than the uncorrected number (Menanteau et al. 2010a). We report  $N_{galc}$  for the remainder of this work.

### 4.3. Recovery of the Brightest Cluster Galaxies

We have designed our observations to detect BCGs to  $z \sim 1.5$ . To quantify the actual depth of our images, we perform a Monte Carlo simulation by injecting artificial sources and computing their recovery fraction. We create the artificial sources with the MODELING package, part of ASTROPY (The Astropy Collaboration et al. 2013).

Following the procedure given in Menanteau et al. (2010b), the synthetic galaxies are created to have de Vaucouleurs (de Vaucouleurs 1948) profiles and surface brightnesses corresponding to their magnitude and assumed sizes. We inject the artificial galaxies into our science images with similar noise characteristics as their real counterparts.

We generate four rounds of one hundred elliptical galaxies spread randomly across our science imaging. Each round of galaxies are placed at different random positions to suppress abnormally boosted recovery fractions due to source confusion. The artificial galaxies have total fluxes corresponding to apparent magnitudes between 19 mag  $< i < 27$  mag with 0.1 mag spacing.

This is almost directly taken from FM2010 – edit. We use the individual field's completion limit to estimate the redshift to which we can reliably detect massive clusters. For this, we compare the completeness limits of our observations to the expected and observed (i.e., known) apparent magnitudes of galaxies in clusters as a function of redshift. We estimated the expected apparent galaxy  $i$ -band magnitude as a function of redshift using  $L_*$  as defined for the population of red galaxies by Blanton et al. (2003) at  $z = 0.1$  and allowing passive evolution according to a solar metallicity (Bruzual & Charlot 2003)  $\tau = 1.0$  Gyr burst model formed at  $zf = 5$ . We show this in Figure 2 for a range of luminosities ( $L_*$ ,  $0.4L_*$ , and  $4L_*$ ) aimed at representing the cluster members from the faint ones to the BCG.

## 5. RESULTS

In this section, we give the results of our cluster finding. We report high confidence, high richness clusters. A high confidence result consists of a clear BCG and many accompanying satellite galaxies. For the 85 fields observed

with MOSAIC, we observe fifteen high confidence clusters (see Figures 3–6). In the following subsections we present on each of the eight high confidence observations individually, and group the medium and low confidence observations together.

### 5.1. Notes on Specific Clusters

In the following subsections, we note previously known sources by querying the NASA/IPCA Extragalactic Database (NED)<sup>5</sup> and the SIMBAD (Set of Identifications, Measurements, and Bibliography for Astronomical Data) astronomical database<sup>6</sup> (Wenger et al. 2000). We include sources from the NRAO (National Radio Astronomy Observatory) VLA (Very Large Array) Sky Survey (NVSS; Condon et al. 1998), the Röntgensatellit (ROSAT) All-Sky Survey Bright Source Catalog (RASS-BSC; Voges et al. 1999), the ROSAT All-Sky Faint Source Catalog (RASS-FSC; Voges et al. 2000), and the SDSS. We make note of confident associations of X-ray and radio sources with the BCG or other clusters members within 5' of the reported BCG pointing (within 10' for the three low redshift clusters).

#### 5.1.1. PSZ2 G029.66-47.63

This is a rich cluster at  $z_{cl} = 0.34 \pm 0.03$  with 130 members, approximately 5' to the northwest of the *Planck* position. The X-ray source 1RXS J214531.1–214339 is coincident with the BCG. Our data show another, slightly less rich, system (with  $z_{cl} = 0.33 \pm 0.04$  and 76 members) within 0.5' of the *Planck* position. Both systems likely contribute to the *Planck* SZ signal. This is the richest cluster in our sample.

#### 5.1.2. PSZ2 G043.44-41.27

This is the second richest cluster we have found; the system is at  $z_{cl} = 0.41 \pm 0.03$  with 116 members. There are two plausible BCGs with nearly the same photo- $z$ ; the one we select is slightly brighter in the  $i$ -band, yields a slightly higher number of cluster members, and is positionally coincident with the X-ray source 1RXS J213644.4–101904. The other bright galaxy is at R.A.=21:36:38.6, decl.=−10:18:35.7, some 1.3' to the west. This galaxy is associated with a bright radio source (NVSS J213638–101836) with a flux density of  $107.8 \pm 3.3$  mJy at 1.4 GHz that has been classified as a symmetric double (?).

#### 5.1.3. PSZ2 G084.62–15.86

This cluster was previously confirmed by the *Planck* team, where they quote a spectroscopic cluster redshift (from two members) of  $z_{spec} = 0.364$  (Planck Collaboration et al. 2016). This system has at least three bright member galaxies within 2' of the *Planck* position that are plausible BCG candidates. Among these we chose the brightest one in the  $i$ -band (which is about an arcminute south of the *Planck* team's chosen BCG) and recovered the cluster at  $z_{photo} = 0.27 \pm 0.10$  with 18 members. The *Planck* team's selected BCG is associated with a radio source (NVSS J214940+331031) with a flux density of  $19.7 \pm 0.8$  mJy at 1.4 GHz.

<sup>5</sup> <https://ned.ipac.caltech.edu/>

<sup>6</sup> <http://simbad.u-strasbg.fr/simbad/>

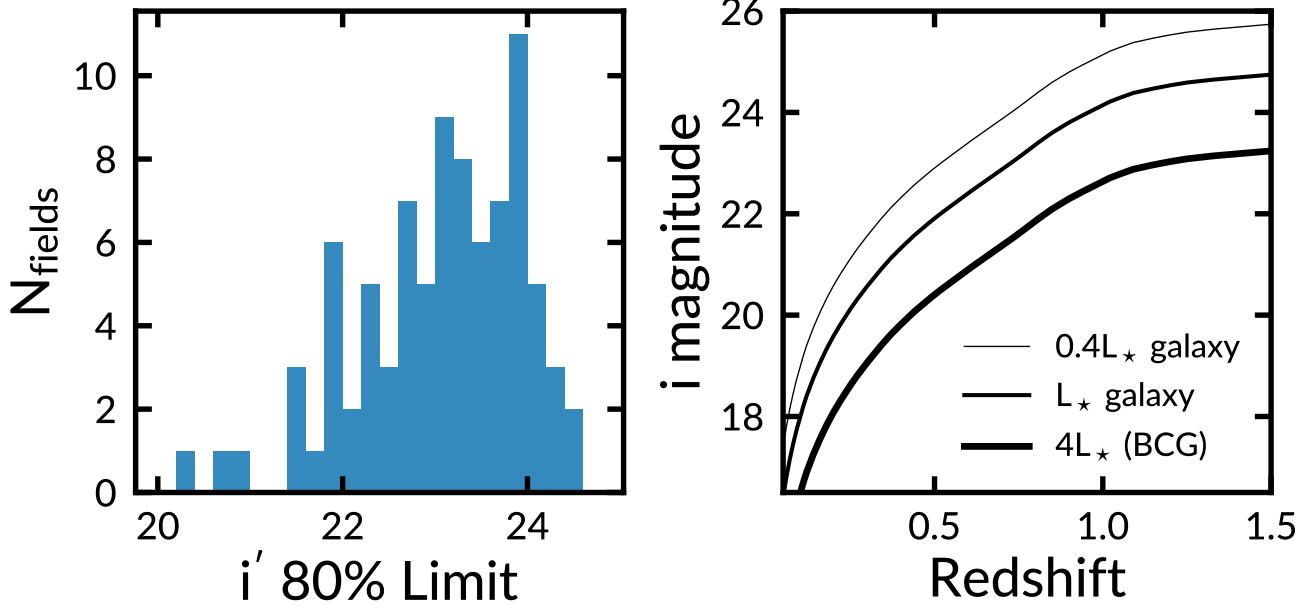


FIG. 2.— *Left:* Histogram of the  $i$ -band magnitude corresponding to 80% completeness in galaxy recovery. When 80% completeness is not achieved we show the limiting magnitude with the highest completeness. *Right:* Observed  $i$ -band magnitudes of  $L_*$ ,  $0.4L_*$ , and  $4L_*$  (BCG) early-type galaxies as a function of redshift. We define an  $LL_*$  galaxy following Blanton et al. (2003) as a population of red galaxies at  $z = 0.1$  and allow it to evolve passively. The left and right panels can be combined to estimate the limiting redshift to which we could identify galaxy clusters.

TABLE 1  
SUMMARY OF CLUSTER FINDING: COLUMN 1: THE PSZ2 CLUSTER NAME; COLUMN 2: PSZ2 SIGNAL-TO-NOISE RATIO; COLUMN 3: PSZ1 ID NUMBER; COLUMN 4: BCG RIGHT ASCENSION IN J2000; COLUMN 5: BCG DECLINATION IN J2000; COLUMN 6: BCG SEPARATION FROM PSZ POSITION IN ARCMINTUES; COLUMN 7: CLUSTER PHOTOMETRIC REDSHIFT WITH  $1-\sigma$  UNCERTAINTY; COLUMN 8: CORRECTED NUMBER OF MEMBER GALAXIES; COLUMN 9: NEW CONFIRMATION?

| Cluster<br>(1)       | SNR<br>(2) | PSZ1 ID<br>(3) | $\alpha$ (J2000)<br>(4) | $\delta$ (J2000)<br>(5) | Sep. ('')<br>(6) | $z_{cl}$<br>(7) | $N_{gal_c}$<br>(8) | New<br>(9) |
|----------------------|------------|----------------|-------------------------|-------------------------|------------------|-----------------|--------------------|------------|
| PSZ2 G029.66 – 47.63 | 5.74       | ...            | 21 : 45 : 29.940        | -21 : 43 : 26.29        | 4.73             | $0.34 \pm 0.03$ | 130                | ✓          |
| PSZ2 G043.44 – 41.27 | 5.55       | ...            | 21 : 36 : 43.728        | -10 : 19 : 02.15        | 1.24             | $0.42 \pm 0.04$ | 116                | ✓          |
| PSZ2 G084.62 – 15.86 | 6.01       | 284            | 21 : 49 : 42.524        | +33 : 09 : 17.29        | 1.14             | $0.27 \pm 0.10$ | 18                 |            |
| PSZ2 G096.43 – 20.89 | 5.81       | ...            | 22 : 48 : 09.402        | +35 : 33 : 49.45        | 0.49             | $0.24 \pm 0.04$ | 54                 | ✓          |
| PSZ2 G098.38 + 77.22 | 5.51       | 346            | 13 : 18 : 08.274        | +38 : 30 : 20.10        | 5.84             | $0.78 \pm 0.05$ | 58                 | ✓          |
| PSZ2 G106.11 + 24.11 | 5.70       | ...            | 19 : 21 : 31.852        | +74 : 33 : 27.17        | 0.51             | $0.14 \pm 0.06$ | 0                  | ✓          |
| PSZ2 G107.83 – 45.45 | 7.09       | ...            | 00 : 07 : 35.605        | +16 : 07 : 02.39        | 0.83             | $0.54 \pm 0.05$ | 30                 | ✓          |
| PSZ2 G120.76 + 44.14 | 5.59       | ...            | 13 : 12 : 53.537        | +72 : 55 : 06.05        | 2.02             | $0.35 \pm 0.03$ | 92                 | ✓          |
| PSZ2 G125.55 + 32.72 | 6.49       | ...            | 11 : 25 : 34.008        | +83 : 58 : 55.75        | 1.44             | $0.21 \pm 0.06$ | 32                 | ✓          |
| PSZ2 G137.24 + 53.93 | 7.87       | ...            | 11 : 40 : 59.525        | +61 : 07 : 07.07        | 4.61             | $0.45 \pm 0.06$ | 40                 | ✓          |
| PSZ2 G173.76 + 22.92 | 5.80       | ...            | 07 : 17 : 26.636        | +44 : 05 : 02.97        | 1.62             | $0.14 \pm 0.04$ | 0                  | ✓          |
| PSZ2 G191.82 – 26.64 | 6.17       | 646            | 04 : 38 : 28.283        | +04 : 37 : 19.91        | 5.18             | $0.18 \pm 0.06$ | 24                 | ✓          |
| PSZ2 G206.45 + 13.89 | 5.90       | 682            | 07 : 29 : 51.234        | +11 : 56 : 31.31        | 1.97             | $0.40 \pm 0.05$ | 73                 |            |
| PSZ2 G224.82 + 13.62 | 5.51       | 752            | 08 : 01 : 41.492        | -04 : 03 : 44.48        | 0.17             | $0.24 \pm 0.04$ | 38                 |            |
| PSZ2 G305.76 + 44.79 | 5.72       | 1070           | 12 : 59 : 53.612        | -18 : 01 : 35.05        | 0.44             | $0.74 \pm 0.08$ | 48                 | ✓          |

#### 5.1.4. PSZ2 G096.43-20.89

The BCG of this cluster is only  $0.5'$  from the *Planck* position. The cluster's redshift is  $z_{cl} = 0.24 \pm 0.04$  with 54 members. An X-ray source (1RXS J224806.6+353230) is  $1.44'$  away from the BCG position. This cluster appears to extend more toward the southwest quadrant in the direction of a catalogued Zwicky cluster (ZwCl 2245.6+3516; Zwicky & Kowal 1968).

#### 5.1.5. PSZ2 G098.38+77.22

This is the highest redshift cluster in the sample at  $z_{cl} = 0.78 \pm 0.05$  with 58 members; it is  $5.8'$  away from the *Planck* position. About  $2'$  east of our chosen BCG there is a luminous ( $i \sim 20$  mag), red galaxy (SDSS J131814.99+383055.8) with a spectroscopic redshift of  $z = 0.726$ . This galaxy has similar colors to the cluster BCG and may be part of the system.

#### 5.1.6. PSZ2 G106.11+24.11

PSZ2 G106.11+24.11 is a low redshift system at  $z_{cl} = 0.14 \pm 0.06$ . The BCG is a large galaxy close to the

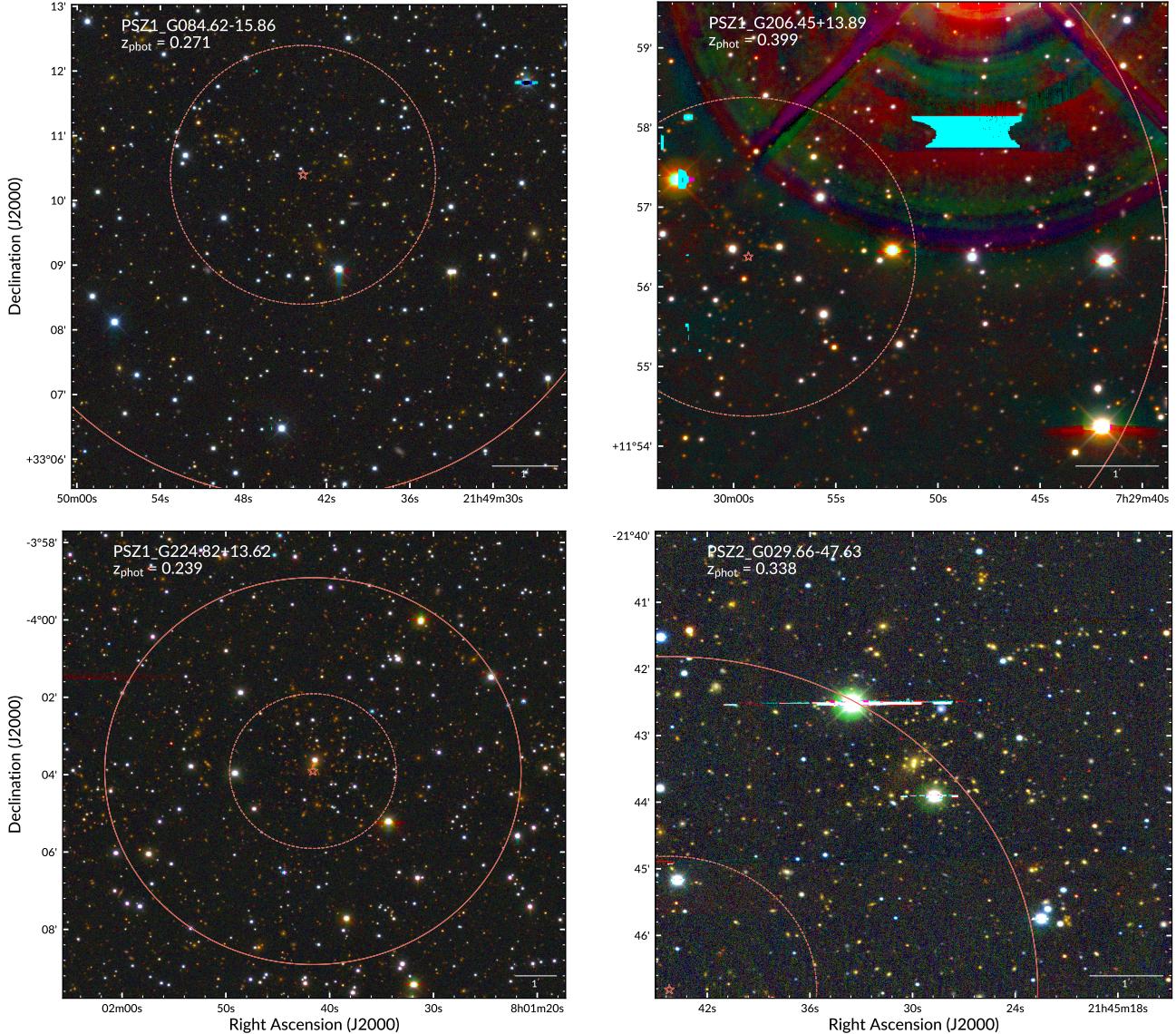


FIG. 3.— RGB (*irg*) color images for four PSZ clusters optically confirmed using the our optical imaging. Each panel is centered on the cluster’s BCG and has a width of 1 mpc at the corresponding cluster’s redshift. The horizontal bar in the lower left of each panel shows the scale of the panel, where north is up and east is to the left. The location of the PSZ detection is denoted by a red star. The dashed and solid concentric circles are 2' and 5' in radius respectively.

*Planck* position. There are 33 members before correcting for background [revise the Table to quote a better estimate of Ngals](#). This object was identified as an X-ray cluster (RXC J1921.3+7433) based on its extent in the RASS (Bohringer et al. 2000), although these authors did not publish a redshift. Using their flux value and our photometric redshift, we estimate the cluster’s X-ray luminosity in the 0.1-2.4 keV band to be  $L_X \sim 2.6 \times 10^{44}$  erg s $^{-1}$ . This luminosity value is broadly consistent with those of other confirmed *Planck* clusters at this redshift range (Planck Collaboration et al. 2015b).

#### 5.1.7. PSZ2 G107.83-45.45

This cluster at  $z_{cl} = 0.54 \pm 0.05$  has 30 members. The BCG (SDSS J000735.62+160701.8) as well as another member (SDSS J000736.15+160508.9) have spectroscopic redshifts from the SDSS of  $z = 0.5673$  and  $z = 0.5667$ , respectively. There are 4 other galaxies with

SDSS spectroscopic redshifts (0.5649, 0.5661, 0.5655, and 0.5625) within 7' (2.7 Mpc) of the BCG.

#### 5.1.8. PSZ2 G120.76+44.14

The BCG we select yields an estimated redshift for the cluster of  $z_{cl} = 0.35 \pm 0.03$  with 92 members; this galaxy has a published spectroscopic redshift of  $z = 0.2959$  (Huchra et al. 1990). We associate the cluster with Abell 1705 and the RASS X-ray source 1RXS J131252.0+725514.

Approximately 2.6' south of the PSZ2 position, there is a luminous red galaxy with a gravitationally lensed arc that is the BCG of a rich system catalogued as WL 1312.5+7252 with a photometric redshift of 0.55 (Dahle et al. 2003). This galaxy is also a radio source (NVSS J131230+725051) with a flux density of  $3.5 \pm 0.5$  mJy (1.4 GHz). Our analysis also yields a rich cluster with this BCG ( $\alpha=13:12:30.9$ ,  $\delta=+72:50:54.2$ ) with  $Ngal_c =$

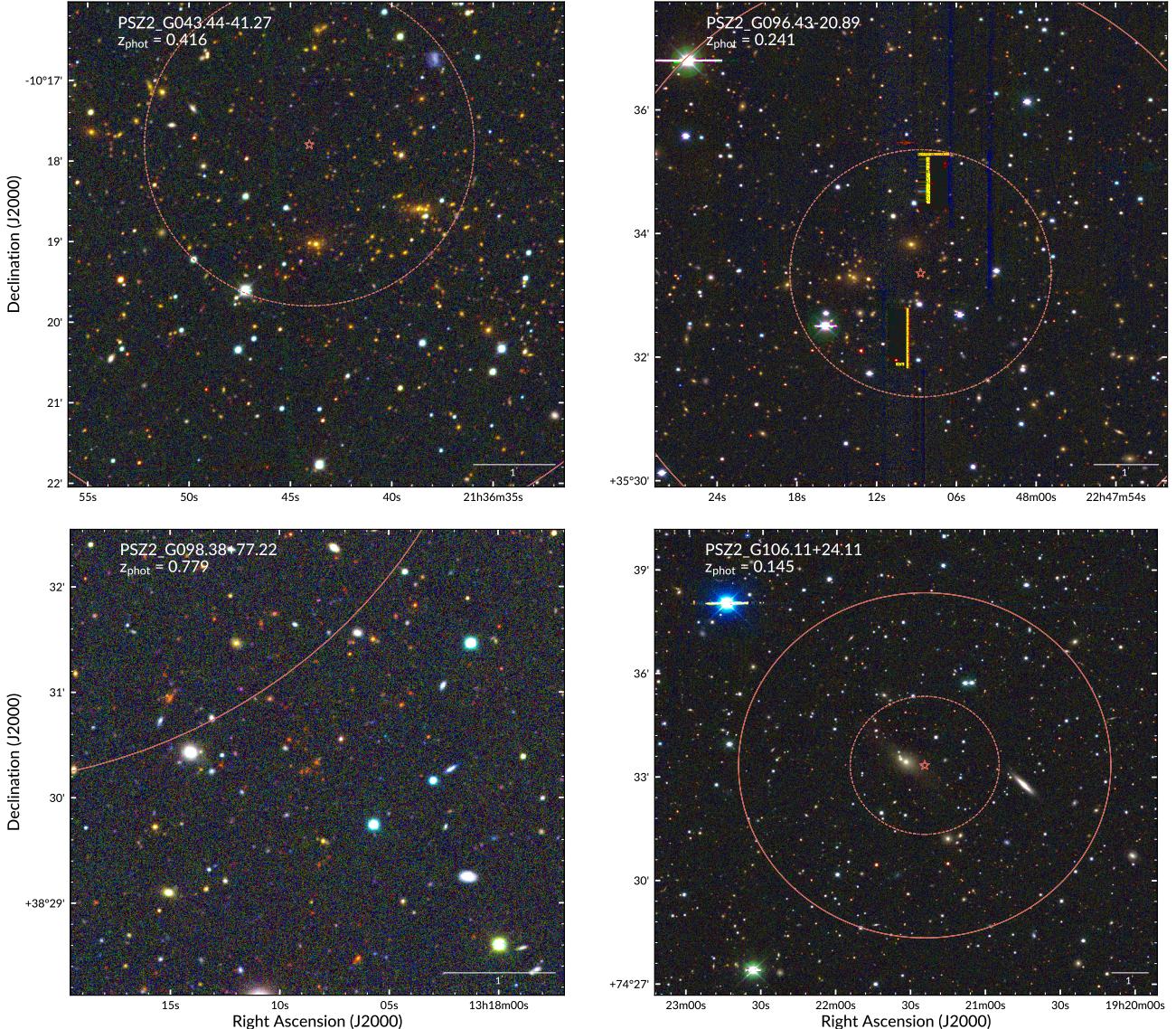


FIG. 4.— RGB (*irg*) color images for four PSZ clusters optically confirmed using the our optical imaging. Each panel is centered on the cluster’s BCG and has a width of 1 mpc at the corresponding cluster’s redshift. The horizontal bar in the lower left of each panel shows the scale of the panel, where north is up and east is to the left. The location of the PSZ detection is denoted by a red star. The dashed and solid concentric circles are 2' and 5' in radius respectively.

89 at  $z = 0.59 \pm 0.05$ .

Given the similar optical richnesses of these two systems it is likely that both contribute to the *Planck* SZ signal. **Or maybe we should say this: On the other hand, given the strong mass sensitivity of SZ selection, perhaps the southern cluster is more likely to be the source of the SZ signal?**

#### 5.1.9. PSZ2 G125.55+32.72

There are two plausible BCGs in this cluster; the one we select yields a redshift of  $z_{cl} = 0.21 \pm 0.06$  and 32 members. The other BCG (R.A.=11:25:46.8, decl.=+83:55:04.4) is about 0.14 mag fainter in the *i*-band and, using it for cluster finding, results on a cluster at  $z_{cl} = 0.20 \pm 0.03$  with 31 members. These two plausible BCGs are separated by  $\sim 4'$ . Both BCGs are radio sources: the northern galaxy (corresponding to our selected BCG) is catalogued as NVSS J112535+835858 with

a 1.4 GHz flux density of  $10.2 \pm 0.9$  mJy; the southern one corresponds to NVSS J112550+835508 with a flux density at the same frequency of  $3.6 \pm 0.6$  mJy. The southern galaxy is about  $1'$  away from the RASS X-ray source 1RXS J112547.3+835559. Both systems should contribute to the *Planck* SZ signal and the quoted richness in Table 1 is almost surely an underestimate of the richness of the combined system.

#### 5.1.10. PSZ2 G137.24+53.93

Here we find a cluster at  $z_{cl} = 0.45 \pm 0.06$  with 40 members. The BCG is a radio source (NVSS J114059+610658) with a flux density of  $10.2 \pm 0.9$  mJy (1.4 GHz) and has a spectroscopic redshift from the SDSS of  $z = 0.4770$ . Another cluster member also has a concordant SDSS spectroscopic redshift of  $z = 0.4697$ . The cluster has been catalogued as WHL J114058.8+610631 with a similar photometric redshift.

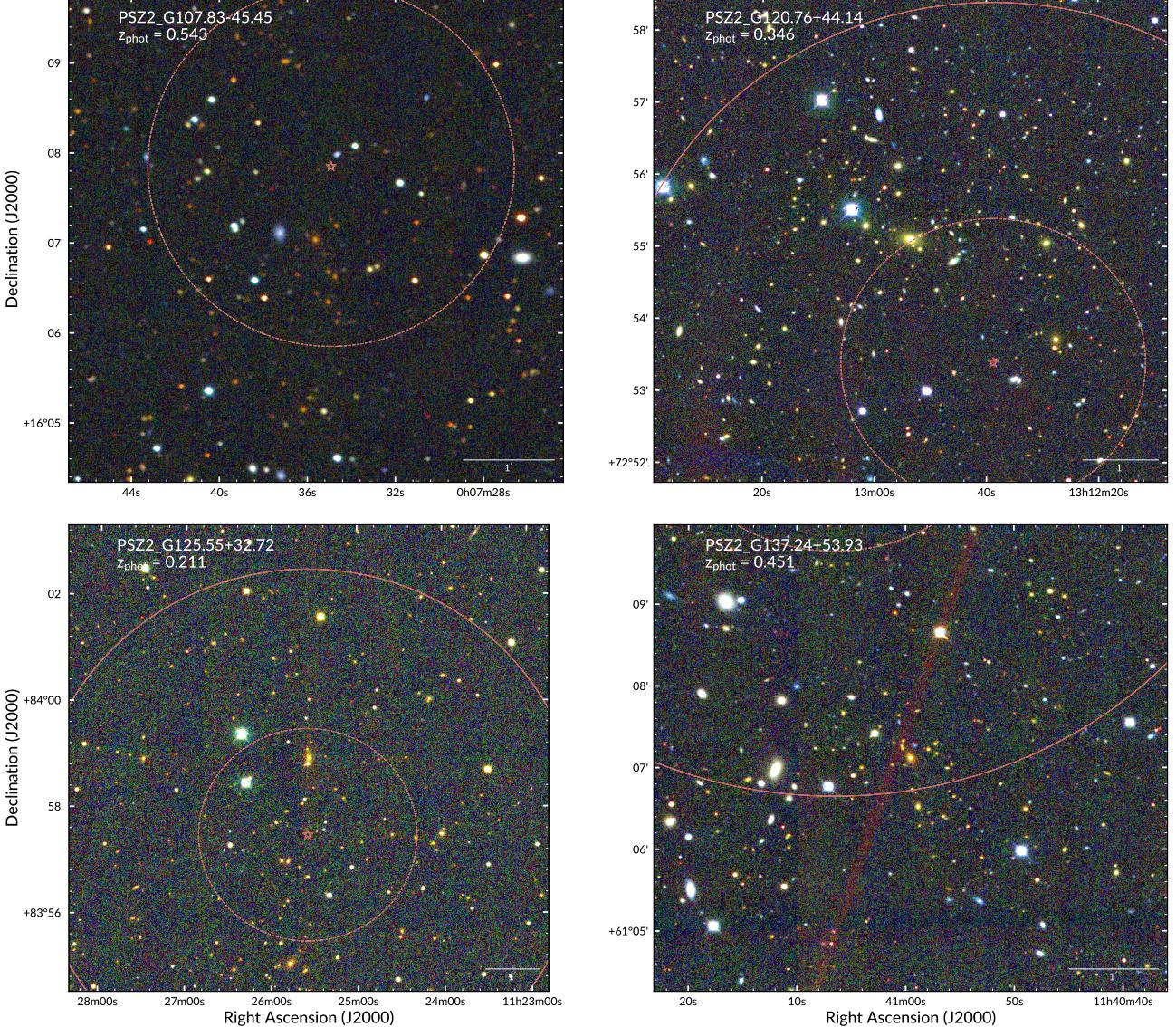


FIG. 5.— RGB (*irg*) color images for four PSZ clusters optically confirmed using the our optical imaging. Each panel is centered on the cluster’s BCG and has a width of 1 mpc at the corresponding cluster’s redshift. The horizontal bar in the lower left of each panel shows the scale of the panel, where north is up and east is to the left. The location of the PSZ detection is denoted by a red star. The dashed and solid concentric circles are 2' and 5' in radius respectively.

#### 5.1.11. PSZ2 G173.76+22.92

This low redshift system, which we find at  $z_{cl} = 0.14 \pm 0.04$ , has a very interesting BCG. It is catalogued in NED as B3 0713+441 and has a spectroscopic redshift of  $z = 0.0652$  (Bauer et al. 2000). It is associated with the RASS X-ray source 1RXS J071726.9+440557 as well as a bright radio source (NVSS J071726+440504) with a flux density of  $220.4 \pm 7.6$  mJy (at 1.4 GHz). Higher resolution images from FIRST reveal that this radio source is double lobed. Some 6' to the east is a catalogued Seyfert 1 galaxy (2MASX J07180060+4405271) with a spectroscopic redshift of  $z = 0.0614$  (Michel & Huchra 1988) and a 1.4 GHz radio flux of  $50.8 \pm 1.6$  mJy (NVSS J071800+440527).

#### 5.1.12. PSZ2 G191.82-26.64

This is another low redshift cluster at  $z_{cl} = 0.18 \pm 0.06$  with 24 members. Two cluster members are associated with radio sources: NVSS J043836+043824 and NVSS J043818+043802 with 1.4 GHz flux densities of  $24.8 \pm 1.2$  mJy and  $22.7 \pm 1.5$  mJy, respectively.

#### 5.1.13. PSZ2 G206.45+13.89

We find a cluster at  $z_{cl} = 0.39 \pm 0.05$  with 73 cluster members. A bright star ( $V = 4.5$  mag; Høg et al. 2000) lies only  $\sim 4.9'$  away from the reported BCG, which prevents accurate photo- $z$  estimates for a significant fraction of the projected area of the cluster. This cluster has been previously confirmed in Barrena et al. (2018) as a rich cluster with a spectroscopic redshift of  $z_{spec} = 0.406$  from 45 members. We confirm the presence of a possible gravitationally-lensed arc  $\sim 13''$  northeast of the BCG.

#### 5.1.14. PSZ2 G224.82+13.62

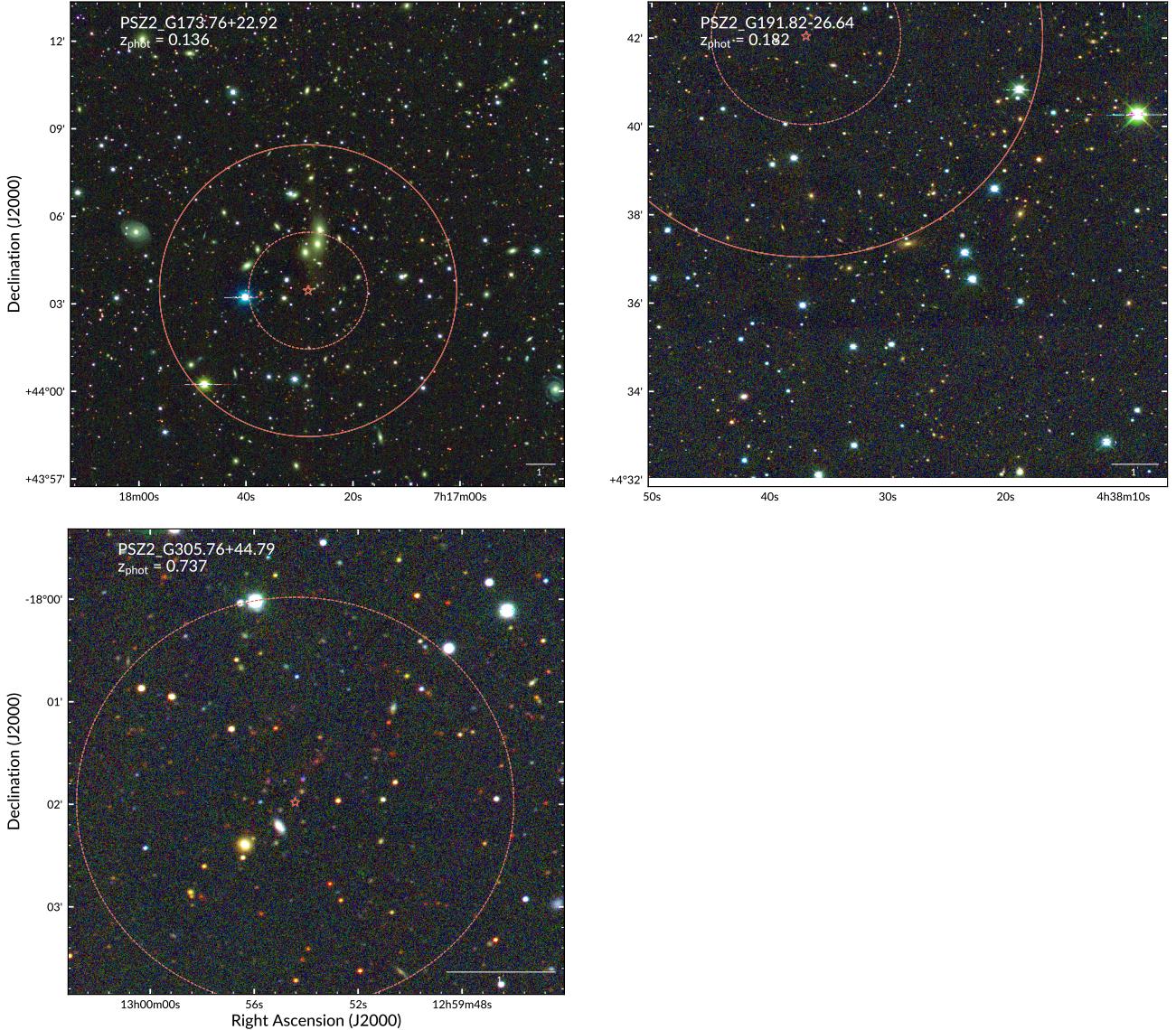


FIG. 6.— RGB (*irg*) color images for four PSZ clusters optically confirmed using the our optical imaging. Each panel is centered on the cluster’s BCG and has a width of 1 mpc at the corresponding cluster’s redshift. The horizontal bar in the lower left of each panel shows the scale of the panel, where north is up and east is to the left. The location of the PSZ detection is denoted by a red star. The dashed and solid concentric circles are 2' and 5' in radius respectively.

The BCG of this system is partially obscured by a nearby star and was not fully deblended in our catalogs. Still we are able to find a rich cluster at  $z_{cl} = 0.23 \pm 0.05$  with 38 members. An interesting aspect of this cluster is that it is positionally coincident with an unidentified X-ray source (2E 0759.2–0355) from the *Einstein Observatory* (Harris et al. 1990). This cluster was confirmed by Barrena et al. (2018) with a spectroscopic redshift of  $z_{spec} = 0.274$  from 28 members.

#### 5.1.15. PSZ2 G305.76+44.79

Finally, PSZ2 G305.76+44.79 is our second highest redshift cluster at  $z_{cl} = 0.74 \pm 0.08$  with 48 members. The BCG is associated with radio source PMN J1259–1801 with a 1.4 GHz flux density of  $42.2 \pm 1.4$  mJy from the NVSS.

## 6. DISCUSSION

In this section, we discuss the results given in the previous section as a whole, and frame those results in the context of the broader PSZ sample. From the 85 fields observed, we identified fifteen rich cluster systems at  $0.1 < z_{cl} < 0.8$ . Because our observations are limited to objects with  $\text{SNR} > 5\sigma$ , we would expect at most one failed detection. This leads to three possible alternatives.

1. The vast majority of clusters in our sample at low- $z$ .
2. The vast majority of the clusters in our sample lie at redshifts beyond our optical detection limits.
3. The vast majority of clusters are obscured by the Milky Way.
4. The vast majority of cluster candidates in our sample are not true  $5\sigma$  detections.

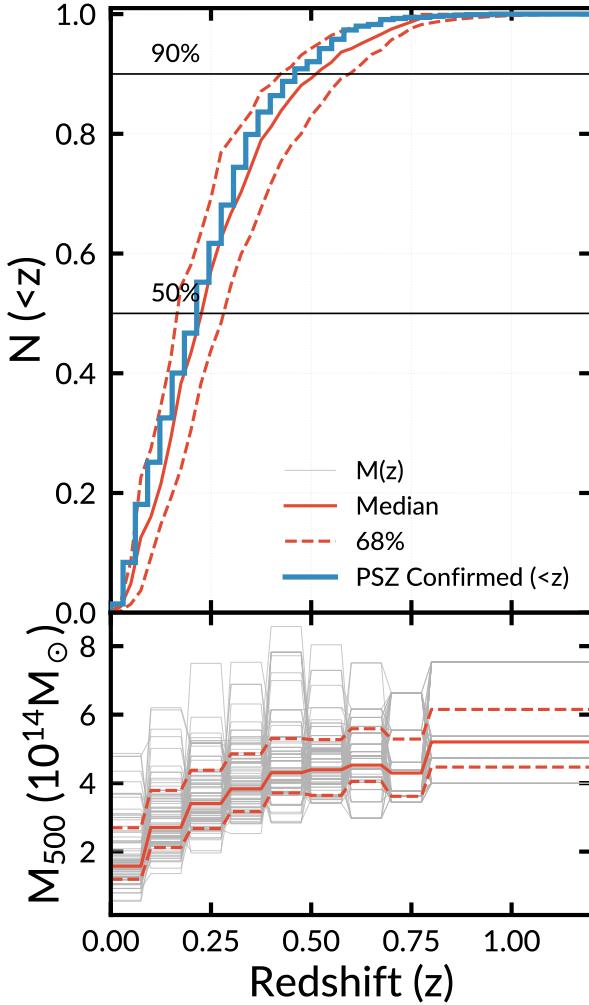


FIG. 7.— Top: The predicted number of clusters as a function of redshift and normalized to unity from the mass function of Tinker et al. (2008). At each redshift the lower mass limit is given by the mass in the bottom panel at the corresponding redshift. The blue PSZ Confirmed curve shows the normalized, cumulative number of confirmed clusters as a function of redshift. For reference we show the 50% and 90% fractional completeness levels. Bottom: Cluster mass as a function of redshift. Redshift bins are the same as Planck Collaboration et al. (2015a) Figure 27. In each panel the solid and dashed orange lines show the median and region enclosing 68% of the data respectively.

**One:** The initial design of our survey uses an  $8' \times 8'$  search window centered on the PSZ position. Low- $z$  clusters could appear as isolated elliptical galaxies and thus would not be classified as clusters. In follow up inspection of the full-sized mosaics, approximately 1 degree $^2$ , reveals no such low- $z$  structures. In addition, we compute the expected fraction of galaxy clusters as a function of redshift from the cluster mass function of Tinker et al. (2008). The cluster mass sensitivity of *Planck* varies with redshift (see Figure 27, Planck Collaboration et al. 2015a). To capture this variability, we draw random (with replacement) samples from the confirmed PSZ2 clusters. Because each cluster is previously confirmed the PSZ2 catalog provides an  $M_{500}$  estimate of the total mass. The bottom panel of Figure 7 shows the cluster mass as a function of redshift, where the gray lines

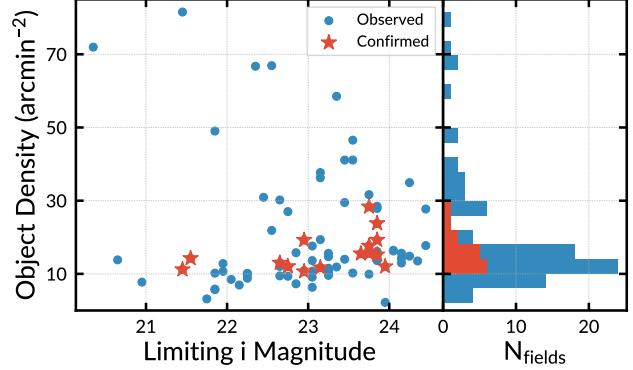


FIG. 8.— Right: Object Density in units of number per arcmin $^{-2}$  as a function of the 80%  $i$ -band limiting magnitude taken from Figure 2. Fields where no cluster is identified are given by blue circles and fields with confirmed clusters are shown by orange stars. Left: Histogram of the object surface density of observed fields.

are individual samples and the orange solid and dashed curves are the median and 68% limits respectively.

The top panel of Figure 7 shows the expected cumulative distribution of galaxy clusters as a function of redshift.

From these numerical simulations we know that a low fraction ( $< 5\%$ ) of all clusters detectable by *Planck* are expected to reside at redshifts less than 0.1. Of the 85 fields observed, we expect approximately four clusters to lie at such low redshifts. We recover three low- $z$  clusters during the course of our analysis.

**Two:** The median limiting  $i$ -band magnitude of our survey is 23 mag, corresponding to a limiting redshift of  $z = 0.72$  for an  $M_\star$  galaxy (see Figure 2). While it is possible that a number of cluster candidates corresponding to real clusters exists above this redshift, it is unlikely that these correspond to a population of the high SNR objects targeted by this survey. Massive clusters, such as those targeted by this survey, are exceedingly rare objects. For example there are only twelve clusters at  $z > 0.72$  in PSZ2, only five of which are above  $z = 0.8$ . Again, Figure 7 allows us to estimate the number of clusters missed with our observing limits. At  $z > 0.7$ , we expect  $\sim 18\%$  of clusters with  $M_{500c} = 3 \times 10^{14} M_\odot$ , corresponding to only fifteen clusters of our 85 observed fields. While high- $z$  objects were expectedly missed by this survey, high- $z$  objects along cannot account for the high number of still unconfirmed clusters. We expect further follow up observations with deep infrared imaging will be required to place further limits on these high- $z$  objects.

**Three:** It is possible that the underlying cause for failed cluster identification is confusion from source crowding in our images. If a cluster lies behind a dense foreground of Milky Way sources, then it could be difficult to visually identify the cluster. To estimate the object surface density in our fields, we sum the number of objects reported by SExtractor and dividing by the sky-area of the image. We explore the possibility that we fail to confirm a cluster in the majority of fields because of foreground contamination in Figure 8. The right panel shows the surface density of objects in the search area as a function of the limiting  $i$ -band magnitude (see also the left panel of Figure 2). The right panel shows the number of

observed fields with the corresponding object density. In both panels, the fields where we do not identify a cluster is shown in blue, and the successfully identified clusters are in orange.

Our identified clusters fall in a relatively narrow range of object densities,  $10 - 30$  Objects arcmin $^{-2}$ . Roughly two thirds of our fields fall within this range. Of the 31 objects outside our identification range, 15 fields have higher densities. If all fifteen fields contain an obscured cluster we could still not account for the vast majority of fields where we did not identify a cluster, 51 fields or  $\sim 60\%$  of our observed sample.

Perhaps more interesting are the sixteen fields where the object surface density is lower than our identification range. These are fields where it should be relatively easy to identify a massive cluster should it exist. Of course, the presence of a massive cluster could raise the object density above the lower identification threshold. These fields, especially the fields with deep optical limiting magnitudes are prime candidates for the potential high- $z$  clusters lurking beyond our reach.

**Four:** Perhaps the most plausible explanation is that no optical counterpart to the PSZ detection exists. Previous works (e.g., Barrena et al. 2018) find a similar fraction of non-detections, and consider both high noise values in the *Planck*  $Y_{500}$  maps (Planck Collaboration et al. 2014a) along with contamination from foreground radio sources. As part of our NED search, we find that approximately 75% of confirmed PSZ sources have a NVSS radio source (39.6 mJy average flux) within 5' of the PSZ position. unconfirmed PSZ sources show slightly fewer sources with approximately 55% having a NVSS radio source (25.1 mJy average flux) within 5'.

5 $\sigma$  corresponds to 99.9999426697% purity or 1/1744278 will be false.

## 7. SUMMARY

following FM2010

In this work, we report on our analysis of seventeen nights of observing spread over three years (2014–2017). We utilize and independently develop a pipeline to process the *griz* imaging taken with both the MOSAIC 1.1 and MOSAIC 2 imagers on the KPNO Mayall-4m telescope. We present the first results from the complete data set of 85 fields, fifteen rich galaxy clusters of which twelve were previously unknown.

The newly discovered clusters range in photometric redshift between  $0.13 < z < 0.74$ . The upper redshift limit is due to the depth of imaging restricting us to clusters at  $z < 0.8$ . This prevents us from finding the most interesting rich clusters at very high redshifts.

A large motivation for this work has come from the recent successes of other SZ follow up programs (e.g., Bayliss et al. 2016; Sifón et al. 2016; Kirk et al. 2015). We

present this sample of clusters to aid in the confirmation of the PSZ sources, and to potentially reveal clusters with interesting astrophysical properties. The PSZ sources which remain unconfirmed have the potential to be the most interesting.

In future work, we will present the properties of lower richness clusters and small groups of galaxies in addition to multi-wavelength studies using the clusters detected as part of this survey.

## ACKNOWLEDGEMENTS

This work was supported by NASA Astrophysics Data Analysis grant number NNX14AF73G and NSF Astronomy and Astrophysics Research Program award number 1615657. This research made use of several open source packages: APLPY, an open-source plotting package for Python hosted at <http://aplpy.github.com>; the IPYTHON package (Perez & Granger 2007); MATPLOTLIB, a Python library for publication quality graphics (Hunter 2007) and ASTROPY, a community developed core Python package for Astronomy (The Astropy Collaboration et al. 2013). IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy under cooperative agreement with the National Science Foundation (Tody 1993). PYRAF is a product of the Space Telescope Science Institute, which is operated by AURA for NASA. Funding for the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the U.S. Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS Web Site is <http://www.sdss.org/>. The SDSS is managed by the Astrophysical Research Consortium for the Participating Institutions. This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<https://www.cosmos.esa.int/gaia>), processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the *Gaia* Multilateral Agreement. This research has made use of the VizieR catalogue access tool, CDS, Strasbourg, France. The original description of the VizieR service was published in Ochsenbein et al. (2000). This research has made use of the SVO Filter Profile Service (<http://svo2.cab.inta-csic.es/theory/fps/>) supported from the Spanish MINECO through grant AyA2014-55216. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

## REFERENCES

- Albareti, F. D., Prieto, C. A., Almeida, A., et al. 2017, The Astrophysical Journal Supplement Series, 233, 25
- Amodeo, S., Mei, S., Stanford, S. A., et al. 2018, The Astrophysical Journal, 853, 36
- Barrena, R., Streblyanska, A., Ferragamo, A., et al. 2018, eprint arXiv:1803.05764, arXiv:1803.05764
- Bauer, F. E., Condon, J. J., Thuan, T. X., & Broderick, J. J. 2000, The Astrophysical Journal Supplement Series, 129, 547
- Bayliss, M. B., Ruel, J., Stubbs, C. W., et al. 2016, The Astrophysical Journal Supplement Series, 227, 3
- Benitez, N. 2000, The Astrophysical Journal, 536, 571
- Bertin, E. 2006, in Astronomical Society of the Pacific Conference Series, Vol. 351, Astronomical Data Analysis Software and Systems XV, ed. C. Gabriel, C. Arviset, D. Ponz, & S. Enrique, 112

- Bertin, E., & Arnouts, S. 1996, *Astronomy and Astrophysics Supplement Series*, 117, 393
- Bertin, E., & Emmanuel. 2011, *Astrophysics Source Code Library*, record ascl:1110.006
- Bertin, E., Mellier, Y., Radovich, M., et al. 2002, in *Astronomical Society of the Pacific Conference Series*, Vol. 281, *Astronomical Data Analysis Software and Systems XI*, ed. D. Bohlender, D. Durand, & T. Handley, 228
- Blanton, M. R., Hogg, D. W., Bahcall, N. A., et al. 2003, *The Astrophysical Journal*, 592, 819
- Bohringer, H., Voges, W., Huchra, J. P., et al. 2000, *The Astrophysical Journal Supplement Series*, 129, 435
- Bruzual, G., & Charlot, S. 2003, *Monthly Notices of the Royal Astronomical Society*, 344, 1000
- Burenin, R. A. 2017, *Astronomy Letters*, 43, 507
- Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *Publications of the Astronomical Society of the Pacific*, 123, 568
- Chabrier, G. 2003, *Publications of the Astronomical Society of the Pacific*, 115, 763
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, eprint arXiv:1612.05560, arXiv:1612.05560
- Coe, D., Benitez, N., Sanchez, S. F., et al. 2006, *The Astronomical Journal*, 132, 926
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *The Astronomical Journal*, 115, 1693
- Dahle, H., Pedersen, K., Lilje, P. B., Maddox, S. J., & Kaiser, N. 2003, *The Astrophysical Journal*, 591, 662
- Dahlen, T., Mobasher, B., Faber, S. M., et al. 2013, *The Astrophysical Journal*, 775, 93
- de Vaucouleurs, G. 1948, *Annales d'Astrophysique*, 11, 247
- Flewelling, H. A., Magnier, E. A., Chambers, K. C., et al. 2016, eprint arXiv:1612.05243, arXiv:1612.05243
- Gaia Collaboration, G., Brown, A. G. A., Vallenari, A., et al. 2016a, *Astronomy & Astrophysics*, Volume 595, id.A2, 23 pp., 595, arXiv:1609.04172
- Gaia Collaboration, G., Prusti, T., de Bruijne, J. H. J., et al. 2016b, *Astronomy & Astrophysics*, Volume 595, id.A1, 36 pp., 595, arXiv:1609.04153
- Harris, D. E., Forman, W., Gioia, I. M., et al. 1990, *The Einstein Observatory catalog of IPC X ray sources. Volume 4E: Right ascension range 08h 00m to 11h 59m*, by D.E. Harris et al. Smithsonian Astrophysical Observatory, 1990, 4
- Hasselfield, M., Hilton, M., Marriage, T. A., et al. 2013, *Journal of Cosmology and Astroparticle Physics*, 2013, 008
- Høg, E., Fabricius, C., Makarov, V. V., et al. 2000, *Astronomy and Astrophysics*, 355, L27
- Huchra, J. P., Geller, M. J., Henry, J. P., & Postman, M. 1990, *The Astrophysical Journal*, 365, 66
- Hunter, J. D. 2007, *Computing in Science & Engineering*, 9, 90
- Ilbert, O., Capak, P., Salvato, M., et al. 2009, *The Astrophysical Journal*, 690, 1236
- Jenkins, A., Frenk, C. S., White, S. D. M., et al. 2001, *Monthly Notices of the Royal Astronomical Society*, 321, 372
- Kirk, B., Hilton, M., Cress, C., et al. 2015, *Monthly Notices of the Royal Astronomical Society*, 449, 4010
- Komatsu, E., Smith, K. M., Dunkley, J., et al. 2011, *The Astrophysical Journal Supplement Series*, 192, 18
- Liu, J., Hennig, C., Desai, S., et al. 2015, *Monthly Notices of the Royal Astronomical Society*, 449, 3370
- Menanteau, F., Hughes, J. P., Jimenez, R., et al. 2009, *The Astrophysical Journal*, 698, 1221
- Menanteau, F., Hughes, J. P., Barrientos, L. F., et al. 2010a, *The Astrophysical Journal Supplement Series*, 191, 340
- Menanteau, F., González, J., Juin, J.-B., et al. 2010b, *The Astrophysical Journal*, 723, 1523
- Menanteau, F., Sifón, C., Barrientos, L. F., et al. 2013, *The Astrophysical Journal*, 765, 67
- Michel, A., & Huchra, J. 1988, *Publications of the Astronomical Society of the Pacific*, 100, 1423
- Molino, A., Benítez, N., Ascaso, B., et al. 2017, *Monthly Notices of the Royal Astronomical Society*, 470, 95
- Mommert, M. 2017, *Astronomy and Computing*, 18, 47
- Ochsenbein, F., Bauer, P., & Marcout, J. 2000, *Astronomy and Astrophysics Supplement Series*, 143, 23
- Oke, J. B. 1974, *The Astrophysical Journal Supplement Series*, 27, 21
- Perez, F., & Granger, B. E. 2007, *Computing in Science & Engineering*, 9, 21
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2011, *Astronomy & Astrophysics*, 536, A1
- . 2014a, *Astronomy & Astrophysics*, 571, A20
- . 2014b, *Astronomy & Astrophysics*, 571, A29
- . 2015a, *Astronomy & Astrophysics*, 594, A27
- . 2015b, *Astronomy & Astrophysics*, 582, A29
- . 2016, *Astronomy & Astrophysics*, 586, A139
- Sifón, C., Battaglia, N., Hasselfield, M., et al. 2016, *Monthly Notices of the Royal Astronomical Society*, 461, 248
- Streblyanska, A., Barrena, R., Rubino-Martin, J. A., et al. 2018, eprint arXiv:1804.01356, arXiv:1804.01356
- Sunyaev, R. A., & Zeldovich, Y. B. 1972, *Comments on Astrophysics and Space Physics*, 4
- Swetz, D. S., Ade, P. A. R., Amiri, M., et al. 2011, *The Astrophysical Journal Supplement Series*, 194, 41
- Tauber, J. A., Mandolesi, N., Puget, J.-L., et al. 2010, *Astronomy and Astrophysics*, 520, A1
- The Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *Astronomy & Astrophysics*, 558, A33
- Tinker, J., Kravtsov, A. V., Klypin, A., et al. 2008, *The Astrophysical Journal*, 688, 709
- Tody, D. 1993, *Astronomical Data Analysis Software and Systems II*, 52
- Valdes, F. G., & Swaters, R. A. 2007, in *Astronomical Society of the Pacific Conference Series*, Vol. 376, *Astronomical Data Analysis Software and Systems XVI*, ed. R. Shaw, F. Hill, & D. Bell, 273
- Voges, W., Aschenbach, B., Boller, T., et al. 1999, *Astronomy and Astrophysics*, 349, 389
- . 2000, *IAU Circulars*, 7432
- Wenger, M., Ochsenbein, F., Egret, D., et al. 2000, *Astronomy and Astrophysics Supplement Series*, 143, 9
- York, D. G., Adelman, J., Anderson, John E., J., et al. 2000, *The Astronomical Journal*, 120, 1579
- Zwicky, F., & Kowal, C. T. 1968, “Catalogue of Galaxies and of Clusters of Galaxies”, Volume VI (Pasadena, California: California Institute of Technology)