

Dynamics

Relative Motion for Rotating Frame

$$\underline{v} = \frac{d\underline{r}}{dt} = \underline{v}_{o'} + \underline{v}_r + \underline{\omega} \times \underline{r}$$

$$\frac{d\underline{v}}{dt} = \underline{a}_r + 2\underline{\omega} \times \underline{v}_r + \frac{d\underline{\omega}}{dt} \times \underline{r} + \underline{a}_{o'} + \underline{\omega} \times (\underline{\omega} \times \underline{r})$$

Polar Coordinates:

$$\underline{v} = \dot{r}\underline{e}_r + r\dot{\theta}\underline{e}_\theta$$

$$\underline{a} = (\ddot{r} - r\dot{\theta}^2)\underline{e}_r + (r\ddot{\theta} + 2\dot{r}\dot{\theta})\underline{e}_\theta$$

Binet's Formula:

For $\underline{a} \times \underline{r} = 0$ we have $a_\theta = 0 \implies r^2\dot{\theta} = c$ which gives

$$a_r = -\frac{c^2}{r^2} \left[\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right]$$

Gravity

$$m\underline{a} = -\frac{GmM}{r^2}\hat{r} = -\frac{m\mu}{r^2}\hat{r}$$

Total Mechanical Energy:

$$E = \frac{1}{2}mv^2 - \frac{m\mu}{r}$$

Angular Momentum:

$$\underline{L} = \underline{r} \times m\underline{v}$$

$$|\underline{L}| = mr^2\dot{\theta} \implies \text{conserved}$$

Keplerian Motion

Two-Body Dynamics

Gravitational Forces:

$$m_i\ddot{\underline{r}}_i = \pm \frac{Gm_im_j}{r^2}\hat{r}$$

Centre-of-Mass:

$$\underline{r}_C = \frac{m_1\underline{r}_1 + m_2\underline{r}_2}{m_1 + m_2}$$

Sum of forces acting on centre of mass = 0 $\implies \underline{r}_C = \underline{r}_0 + \underline{r}'t$

Approximate $\mu \approx Gm_1$ gives:

$$\ddot{\underline{r}} = -\frac{\mu}{r^2}\hat{r}$$

Total energy is conserved, take $\ddot{\underline{r}} \cdot \underline{r}$ to prove $(d/dt)E = 0$, since:

$$\frac{d}{dt} \left(\frac{1}{2}v^2 \right) = \ddot{\underline{r}} \cdot \underline{r} = \frac{d}{dt} \left(\frac{G(m_1 + m_2)}{r} \right)$$

Equations of Motion

Use chain rule and $r^2\dot{\theta} = l$ to get:

$$\frac{dr}{dt} = \frac{l}{r^2} \frac{dr}{d\theta}$$

Change variable to $u = 1/r$:

$$\frac{dr}{dt} = -l \frac{du}{d\theta}$$

Gives:

$$\frac{d^2r}{dt^2} = -l^2u^2 \frac{d^2u}{d\theta^2} \quad \text{and} \quad -r\dot{\theta}^2 = -l^2u^3$$

Equations of motion in polar coordinates are:

$$\ddot{r} - r\dot{\theta}^2 = -\frac{\mu}{r^2} \quad \text{and} \quad r\ddot{\theta} + 2\dot{r}\dot{\theta} = \frac{1}{r} \frac{d}{dt}(r^2\dot{\theta}) = 0$$

Transforming for u :

$$-l^2u^2 \frac{d^2u}{d\theta^2} - l^2u^3 = -\mu u^2$$

$$\frac{d^2u}{d\theta^2} + u = \frac{\mu}{l^2}$$

Finally change to $z = u - \mu/l^2$ to get:

$$\frac{d^2z}{d\theta^2} + z = 0 \implies z = A \cos(\theta - \theta_0)$$

Which gives polar equation of conic section for r :

$$r(\theta) = \frac{l^2/\mu}{1 + \frac{Al^2}{\mu} \cos(\theta)}$$

The energy equation is :

$$E = \frac{1}{2}m(\dot{r}^2 + r^2\dot{\theta}^2) - U(r)$$

Which can be solved by integrating w.r.t r , changing variable from t to θ , and substituting $U(r) = m\mu/r$. This gives:

$$r = \frac{p}{\pm 1 + e \cos(\theta - \theta_0)}$$

where

$$p = \frac{c^2}{|\mu|} = a(1 - e^2) \quad \text{and} \quad e = \sqrt{\frac{2Ec^2}{m\mu^2} + 1}$$

so r describes an ellipse with semimajor axis a and eccentricity e , pericentre at $\theta = 0$ and apocentre at $\theta = \pi$:

$$r = \frac{a(1 - e^2)}{1 + e \cos(\theta - \theta_0)}$$

Orbital Parameters

$$a = \frac{r_p + r_a}{2} \quad \text{and} \quad e = \frac{r_a - r_p}{r_a + r_p}$$

so

$$r_p = a(1 - e) \quad \text{and} \quad r_a = a(1 + e)$$

Comparing equations for r :

$$l^2 = \mu a(1 - e^2)$$

$$v_p = \sqrt{\frac{\mu(1+e)}{a(1-e)}} \quad \text{and} \quad v_a = \sqrt{\frac{\mu(1-e)}{a(1+e)}}$$

Conservation of energy:

$$E = -\frac{\mu}{2a} \implies v^2 = \mu \left(\frac{2}{r} - \frac{1}{a} \right)$$

Kepler's Laws

1st Law: Orbits are elliptical with attracting body at focus of ellipse

2nd Law: Radial vector sweeps out equal areas in equal time

$$\implies T = 2\pi \sqrt{\frac{a^3}{\mu}}$$

3rd Law: Orbital period squared \propto semi-major axis cubed:

Kepler's Equation

Mean motion

$$n = \sqrt{\frac{\mu}{a^3}}$$

Mean anomaly

$$M = n(t - \tau)$$

Kepler's equation is

$$M = E - e \sin E$$

and use trig identity to get θ from E

$$\tan\left(\frac{\theta}{2}\right) = \sqrt{\frac{1+e}{1-e}} \tan\left(\frac{E}{2}\right)$$

This must be solved iteratively, start with guess $E_0 = M$, then find correction

$$\Delta E = \frac{M - E_0 + e \sin E_0}{1 - e \cos E_0}$$

Then repeat, using $E_i = E_{i-1} + \Delta E_{i-1}$ until $\Delta E_i \approx 0$

Position and Velocity from Orbital Elements

From orbit equation

$$\dot{r} = \frac{a(1-e^2)}{(1+e \cos \theta)^2} e \sin \theta \dot{\theta} = r \frac{e \sin \theta}{1+e \cos \theta} \dot{\theta}$$

Flight path angle

$$\tan \gamma = \frac{v_r}{v_\theta} = \frac{\dot{r}}{r\dot{\theta}} = \frac{e \sin \theta}{1+e \cos \theta}$$

To get in plane position and velocity at time t , solve Kepler's equation for τ time of passage from pericentre, then compute orbit radius and thus cartesian positions.

$$x = r \cos \theta \quad \text{and} \quad y = r \sin \theta$$

Velocity components:

$$v_r = \frac{r^2 e \sin \theta}{p} \dot{\theta} \quad \text{and} \quad v_\theta = \frac{l}{r} = r \dot{\theta}$$

Gives cartesian velocities:

$$\dot{x} = -\sqrt{\frac{\mu}{p}} \sin \theta \quad \text{and} \quad \dot{y} = \sqrt{\frac{\mu}{p}} (e + \cos \theta)$$

Orbit Transfers

Hohmann Transfer

Minimum energy transfer between two circular, coplanar orbits is an ellipse with semimajor axis $a = \frac{1}{2}r_1 + r_2$. The change in velocities required are the difference between the velocities of the transfer ellipse at pericentre and apocentre, and the velocities of the circular orbits at r_1 and r_2

$$\Delta v_1 = \sqrt{\frac{2\mu}{r_1} - \frac{2\mu}{r_1 + r_2}} - \sqrt{\frac{\mu}{r_1}}$$

$$\Delta v_2 = \sqrt{\frac{\mu}{r_2}} - \sqrt{\frac{2\mu}{r_2} - \frac{2\mu}{r_1 + r_2}}$$

Parabolic Transfer

For a parabolic transfer where the orbiting body escapes circular orbit at r_1 to infinity, and is then captured from infinity into orbit at r_2

$$\Delta v_1 = (\sqrt{2} - 1) \sqrt{\frac{\mu}{r_1}}$$

$$\Delta v_2 = (\sqrt{2} - 1) \sqrt{\frac{\mu}{r_2}}$$

Single Impulse Manoeuvres

Raise apocentre:

$$\Delta v = \sqrt{-\frac{\mu}{a_2} + \frac{2\mu}{r_p}} - \sqrt{-\frac{\mu}{a_1} + \frac{2\mu}{r_p}}$$

Change inclination:

$$\Delta v = \sqrt{v_1^2 + v_2^2 - 2v_1 v_2 \cos \Delta i}$$

Multiple impulse manoeuvres can be used to reduce required delta v, by first raising apocentre and performing multiple burns.

Patched Conic

In Earth-centred frame, velocity to escape is:

$$\Delta v_{esc} = \sqrt{\frac{2\mu_{Earth}}{r_p} + v_{\infty,1}^2} - \sqrt{\frac{2\mu_{Earth}}{r_p} - \frac{\mu_{Earth}}{a_p}}$$

In sun centre frame, $v_{\infty,1}$ must match the first Δv for transfer. E.g. Δv_1 of Hohmann transfer, and similar for capture.

Rocket Equation

Specific impulse

$$I_{sp} = \frac{F}{\dot{m}g_0}$$

Effective exhaust velocity

$$c = I_{sp}g_0$$

Rocket equation

$$(m - dm)dv = -dmv_e$$

$$\Delta v = v_e \ln \frac{m_i}{m_f}$$

$$m_f = m_i e^{-\frac{\Delta v}{v_e}}$$

Aerobraking

Perturbations

N-Body Dynamics