

1 Nonlinear Evolution of Massive Scalar Fields in Anti-de Sitter Spacetime

As mentioned in §??, within the space of initial data for massive scalar fields in AdS_{d+1} there are islands of stability for Gaussian momentum profiles where collapse does not occur for sufficiently small perturbations. Examining the dependence of the end state (stable, unstable, or otherwise) on the initial conditions has been the goal of previous works, such as [?, ?] and others. However, the limit of small but non-perturbative amplitudes requires significant computing resources and has only recently become computationally accessible. The goal of the work presented here is to leverage the computing resources available through Westgrid and Compute Canada to examine the broadest possible range of initial parameters.

1.1 Contributions of Authors

The research covered in this work built upon on the numerical solving methods first utilized as part of [?, ?] to examine nonlinear instabilities in the full Einstein/Klein-Gordon system of massive scalars in AdS_{d+1} . The goal of this work was to expand the space of initial data being considered and thereby chart the islands of stability. Previous work had identified these islands and commented on the transition regions between stable and unstable configurations, but had avoided the computationally costly simulations required for cases on the “shoreline.”

My role involved running multiple, simultaneous simulations over different heterogeneous computing clusters through the Westgrid network of the Compute Canada consortium. I was responsible for

roughly 184 core years¹ worth of simulations over the course of this work.

To perform the data analysis required, I wrote new plotting programs that interfaced with existing data types while also providing extensions specifically for this project. For example, code for plotting the horizon formation time t_H against the amplitude of the perturbation ϵ was rewritten in python with a range of fitting options for critical data above a specified t_{fit} (see plots and insets in figure 1.4 along with fitting parameters in table 1.1). I also programmed and performed all of the convergences tests included in appendix 1.A that verified the evolution of the data. Of particular importance was discerning the reliability of the solutions for irregular data. In figures 1.A.4 and 1.A.5, the order of convergence Q is calculated for irregular data. These tests help to validate important observations of chaotic evolution, even for massless scalars.

As is common for these types of projects, all members of the collaboration were equally involved in the interpretation of the data, as well as the late stages of editing. Authors are listed alphabetically and it is understood that all members contribute equally to the publication.

¹From the [Compute Canada](#) website: “a core year is the equivalent of running computations on a CPU core constantly for a period of one year.”

Phase Diagram of Stability for Massive Scalars in Anti-de Sitter Spacetime

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We diagram the behavior of 5-dimensional anti-de Sitter spacetime against horizon formation in the gravitational collapse of a scalar field, treating the scalar field mass and width of initial data as free parameters, which we call the stability phase diagram. We find that the class of stable initial data becomes larger and shifts to smaller widths as the field mass increases. In addition to classifying initial data as stable or unstable, we identify two other classes based on nonperturbative behavior. The class of metastable initial data forms a horizon over longer time scales than suggested by the

lowest order perturbation theory at computationally accessible amplitudes, and irregular initial data can exhibit non-monotonic and possibly chaotic behavior in the horizon formation times. Our results include evidence for chaotic behavior even in the collapse of a massless scalar field.

1.2 Introduction

Through the anti-de Sitter spacetime (AdS)/conformal field theory (CFT) correspondence, string theory on $\text{AdS}_5 \times X^5$ is dual to a large N conformal field theory in four spacetime dimensions ($\mathbb{R} \times S^3$ when considering global AdS_5). The simplest time-dependent system to study in this context is the gravitational dynamics of a real scalar field with spherical symmetry, corresponding to the time dependence of the expectation value of the zero mode of a single trace operator in the gauge theory. Starting with the pioneering work of [?, ?, ?, ?], numerical studies have suggested that these dynamics may in fact be generically unstable toward formation of (asymptotically) AdS_{d+1} black holes even for arbitrarily small amplitudes. While perhaps surprising compared to intuition from gravitational collapse in asymptotically flat spacetimes, the dual picture of thermalization of small energies in a compact space is more expected. In terms of the scalar eigenmodes on a fixed AdS background, the instability is a cascade of energy to higher frequency modes and shorter length scales (weak turbulence), which eventually concentrates energy within its Schwarzschild radius. In a naive perturbation theory, this is evident through secular growth terms.

However, some initial scalar field profiles lead to quasi-periodic evolution (at least on the time scales accessible via numerical studies) at small but finite amplitudes; even early work [?, ?] noted that it is possible to remove the secular growth terms in the evolution of a single perturbative eigenmode. A more sophisticated perturbation theory [?, ?, ?, ?, ?, ?, ?, ?, ?, ?, ?] supports a broader class of quasi-periodic solutions that can contain non-negligible contributions from many modes, and other stable solutions orbit the basic quasi-periodic solutions [?]. Stable solutions exhibit inverse cascades of energy from higher frequency to lower frequency modes due to conservation laws following from the high symmetry of AdS (integrability of the dual CFT). Stable behavior also appears in the full

non-perturbative dynamics for initial profiles with widths near the AdS length scale [?, ?, ?]; however, analyses of the perturbative and full dynamics in the literature have not always been in agreement at fixed small amplitudes. For example, some perturbatively stable evolutions at finite amplitude actually form black holes in numerical evaluation of the full dynamics [?, ?, ?]. Understanding the breakdown of the approximations used in the perturbative theory, as well as its region of validity, is an active and important area of research [?, ?, ?, ?, ?].

Ultimately, the main goal of this line of inquiry is to determine whether stability or instability to black hole formation (or both) is generic on the space of initial data, so the extent of the “islands of stability” around single-mode or other quasi-periodic solutions and how it varies with parameters of the physics on AdS are key questions of interest. The biggest changes occur in theories with a mass gap in the black hole spectrum, such as AdS_3 and Einstein-Gauss-Bonnet gravity in AdS_5 , which cannot form horizons at small amplitudes. While small-amplitude evolution in AdS_3 appears to be quasi-periodic [?, ?], there is some evidence to point toward late-time formation of a naked singularity in AdS_5 Einstein-Gauss-Bonnet gravity [?, ?] (along with a power law energy spectrum similar to that at horizon formation). Charged scalar and gauge field matter [?] also introduces a qualitative change in that initial data may lead to stable evolution or instability toward either Reissner-Nordström black holes or black holes with scalar hair.

In this paper, we extend the study of massive scalar matter initiated in [?, ?]. Specifically, using numerical evolution of the full gravitational dynamics, we diagram classes of gravitational collapse behavior as a function of scalar field mass and initial scalar profile width, which we call a stability phase diagram in analogy to a phase diagram for phases of matter. This is the first systematic study of behavior for classes of initial data in AdS gravitational collapse using two tuning parameters. By considering the time to horizon formation as a function of the initial profile’s amplitude at finite amplitude, we identify several different classes of behavior and indicate them on the phase diagram. Finally, we analyze and characterize these different behaviors, presenting evidence for chaotic behavior, including the first evidence for chaotic behavior in the horizon formation time of massless scalar collapse, which has no length scale other than the AdS radius. Throughout, we work

in AdS_5 , due to its relevance to strongly coupled gauge theories in four dimensions and because previous literature has indicated massless scalars lead to greater instability than in AdS_4 (the main other case considered), which makes the effects of the scalar field mass more visible.

We note briefly two caveats for the reader. First, horizon formation always takes an infinite amount of time on the AdS conformal boundary due to the usual time dilation effects associated with horizons; this agrees with the understanding of thermalization in the CFT as an asymptotic process. Horizon formation times discussed in this paper correspond to an approximate notion of horizon formation that we will describe below, but alternate measures of thermalization may be of interest. Second, the black holes we discuss are smeared on the compact X^5 dimensions of the gravitational side of the duality, as in most of the literature concerning stability of AdS, and we are particularly interested in small initial amplitudes that lead to black holes small compared to the AdS scale. As described in [?, ?, ?], small black holes in this situation suffer a Gregory-Laflamme-like instability toward localization on X^5 (which may in fact lead to formation of a naked singularity). At the same time, certain light stable solutions for charged scalars (boson stars) are stable against localization on X^5 [?]. We therefore provisionally assume that the onset of the Gregory-Laflamme-like instability occurs only at horizon formation, not at any point of the earlier horizon-free evolution.

The plan of this paper is as follows: in section 1.3, we review the time scales associated with horizon formation with an emphasis on the behavior of massive scalars and briefly discuss our methods. Then, in section 1.4, we present the phase diagram of different stability behaviors, and an attempt at heuristic analytic understanding appears in 1.5. We close with a discussion of our results.

1.3 Review

In this section, we review results on the stability of scalar field initial data as well as our methods (following the discussion of [?]).

1.3.1 Massive scalars, stability, and time scales

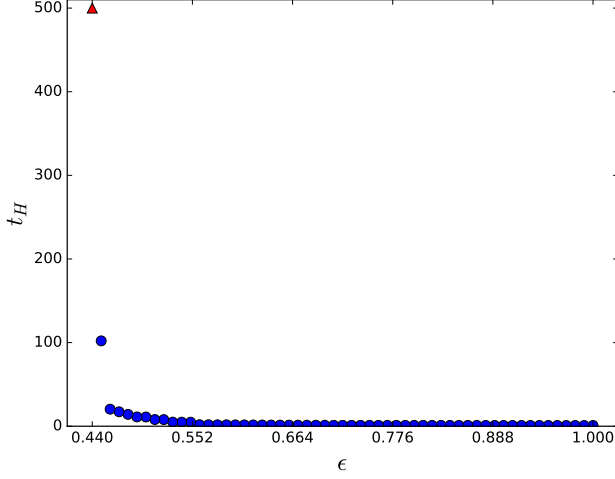
As in most of the literature, we work in Schwarzschild-like coordinates, which have the line element (in asymptotic AdS_{d+1})

$$ds^2 = \frac{1}{\cos^2(x)} \left(-Ae^{-2\delta} dt^2 + A^{-1} dx^2 + \sin^2(x) d\Omega^{d-1} \right) \quad (1.1)$$

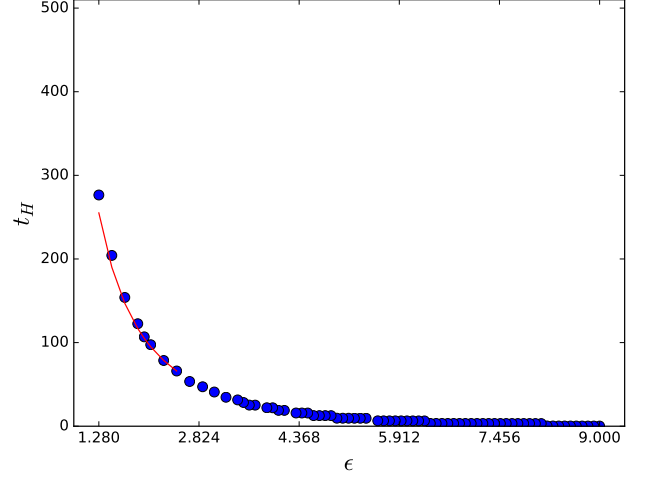
in units of the AdS scale. In these coordinates, a horizon appears at $A(x, t) = 0$, but reaching zero takes an infinite amount of time (measured either in proper time at the origin or in conformal boundary time); following the standard approach, we define a horizon as having formed at the earliest spacetime point (as measured by t) where A drops below a specified threshold defined in §1.3.2 below. Of course, horizon formation represents a coarse-grained description since the pure initial state of the dual CFT cannot actually thermalize; a more precise indicator of approximate thermalization may be the appearance of a power law energy spectrum (exponentially cut off) in the perturbative scalar eigenmodes. This indicator is tightly associated with horizon formation (though see [?, ?] for some counterexamples).

A key feature of any perturbative formulation of the gravitational collapse is that deviations from $A = 1, \delta = 0$ appear at order ϵ^2 , where ϵ is the amplitude of initial data. As a result, horizons can form only after a time $t \sim \epsilon^{-2}$; in the multiscale perturbation theory of [?, ?, ?, ?, ?, ?, ?, ?, ?], there is in fact a scaling symmetry $\epsilon \rightarrow \epsilon', t \rightarrow t(\epsilon/\epsilon')^2$ that enforces the proportionality $t_H \propto \epsilon^{-2}$, where t_H is the (approximate) horizon formation time for unstable initial data at small amplitude.

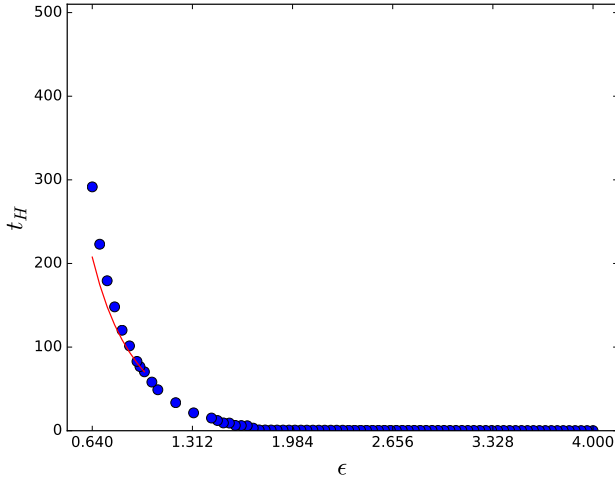
At this point, it is worth making a small clarification. If the collapsing matter takes the form of a well-defined pulse, horizon formation occurs when the pulse nears the origin. For massless matter, that means that the t_H is piecewise continuous as a function of ϵ ; each continuous “step” has approximately constant t_H and is separated from the next step by a time of approximately π , the light crossing time for a round trip from the origin to the boundary of AdS. Massive matter does not reach the boundary, so the steps are not always separated by π , and may in fact not be separated



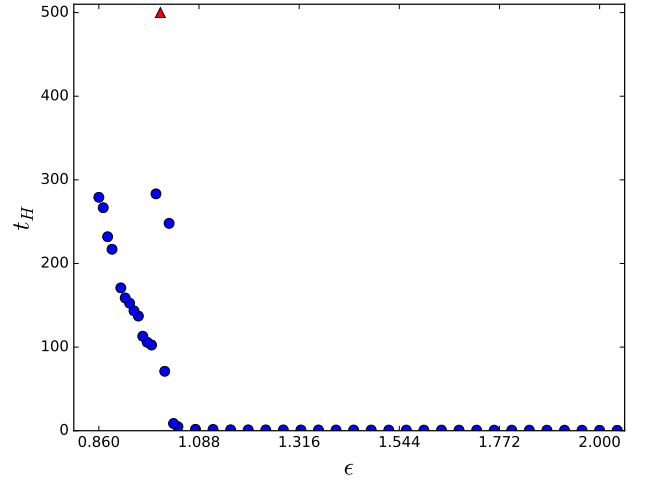
(a) *Stable initial data for $\sigma = 1.5$*



(b) *Unstable initial data for $\sigma = 0.25$*



(c) *Metastable initial data for $\sigma = 0.85$*



(d) *Irregular initial data for $\sigma = 1.1$*

Figure 1.1: *Classes of initial data for massless scalars and initial width σ . Blue dots represent horizon formation; red triangles indicate a lower limit for t_H . Red curves in subfigures 1.1b, 1.1c are $t_H = a\epsilon^{-2} + b$ matched to largest two amplitudes in the curve.*

at all if the pulse spreads out in radius. In any case, though, the width of the steps decreases drastically as amplitude decreases, so it becomes very difficult to find the transition amplitudes numerically. In fact, adjacent amplitudes in a numerical sample are typically multiple steps apart once the evolution is already long, which justifies using the perturbative scaling $t_H \propto \epsilon^{-2}$.

Based on the perturbative scaling relation, initial data can be divided into several classes with

respect to behavior at low amplitudes, as illustrated in figure 1.1 for massless scalars. *Stable* initial data evolves indefinitely without forming a horizon. In practice, we identify this type of behavior in numerical evolutions by noting rapid horizon formation at high amplitude with a vertical asymptote in t_H just above some critical amplitude. In our numerical results, we see a sudden jump at the critical amplitude to evolutions with no horizon formation to a large time t_{lim} , possibly with a small window of amplitudes with large t_H just above the critical amplitude. In a few cases, we have captured a greater portion of the asymptotic region. See figure 1.1a. *Unstable* initial data, in contrast, forms a horizon at all amplitudes following the perturbative scaling relation $t_H \propto \epsilon^{-2}$ as $\epsilon \rightarrow 0$. In our analysis, we will verify this scaling by fitting t_H to a power law as described in section 1.3.2 below; if we limit the fit to smaller values of ϵ , the scaling becomes more accurate. Figure 1.1b shows unstable data. The red curve is of the form $t_H = a\epsilon^{-2} + b$ with a, b determined by matching the curve to the data for the largest two amplitudes with $t_H \geq 60$ (not a best fit); note that the data roughly follows this curve. The categorization of different initial data profiles with similar characteristic widths into stable and unstable is robust for massless and massive scalars [?]; small and large width initial data are unstable, while intermediate widths are stable. One of the major results of this paper is determining how the widths of initial data in these “islands of stability” vary with scalar mass.

A priori, there are other possible types of behavior, at least beyond the first subleading order in perturbation theory, that is, at finite ϵ . *Metastable* initial data collapses with $t_H \propto \epsilon^{-p}$ with $p > 2$ at small but not arbitrarily small amplitudes (or another more rapid growth of t_H with decreasing amplitude). We will find this type of behavior common on the “shoreline” of islands of stability where stable behavior transitions to unstable. As we will discuss further below, metastable behavior may or may not continue as $\epsilon \rightarrow 0$; in principle, as higher order terms in perturbation theory become less important, the behavior may shift to either stable or unstable as described above. In principle, initial data that is stable at third order in perturbation theory but unstable at higher order could have metastable scaling even in the $\epsilon \rightarrow 0$ limit, though our numerical study cannot address this case. We in fact find circumstantial evidence in favor of the different possibilities. In

the case that the $\epsilon \rightarrow 0$ behavior is perturbatively unstable, the perturbative scaling $t_H \propto \epsilon^{-2}$ only appears for larger t_H than the typical unstable case; it may therefore be reasonable for the reader to consider metastable initial data as part of a second order transition between unstable and stable classes of initial data. Figure 1.1c shows metastable initial data that continues to collapse to times $t_H \sim 0.6t_{lim}$ but more slowly than ϵ^{-2} ; note that t_H for collapsed evolutions at small amplitudes lies significantly above the curve $t_H = a\epsilon^{-2} + b$ (which is determined as in figure 1.1b). There was one additional type of behavior identified by [?], which was called “quasi-stable” initial data at the time since the low-amplitude behavior was not yet clear. We find here that these initial data are typically stable at small amplitude but exhibit irregular behavior in t_H as a function of ϵ , so we will denote them as *irregular* initial data; irregular behavior may be strongly non-monotonic or even exhibit some evidence of chaos. Figure 1.1d shows an example of irregular initial data. Later, we will see more striking examples of this behavior for massive scalars.

We emphasize that we are not claiming that metastable or irregular behavior persist to arbitrarily small amplitudes (though a priori metastable behavior could). In that sense, the multiscale perturbation theory suggests that the only two classes of stability behavior are stable and unstable with $t_H \propto \epsilon^{-2}$ scaling as $\epsilon \rightarrow 0$. However, it is also important to understand physics outside the perturbative regime, and classifying the behavior of AdS when higher-order or nonperturbative effects contribute is still of interest. For example, it is clear that metastable initial data (as defined precisely below) does not exhibit perturbatively unstable behavior for t_H values as small as other unstable initial data, even in the cases where it may at all. This may help understanding the breakdown of the multiscale perturbation theory. Similarly, irregular initial data leads to qualitatively different behavior even visually and suggests that nonperturbative dynamics are important. It is in the spirit of looking beyond the multiscale perturbation theory that we call metastable and irregular initial data independent classes of behavior, even if they are not quite on the same standing as perturbatively stable or unstable classes. This paper presents the first systematic mapping of where these distinct behaviors appear.

1.3.2 Methods

For spherically symmetric motion, the Klein-Gordon equation for scalar mass μ can be written in first order form as

$$\phi_{,t} = Ae^{-\delta}\Pi, \quad \Phi_{,t} = (Ae^{-\delta}\Pi)_{,x}, \quad (1.2)$$

$$\Pi_{,t} = \frac{(Ae^{-\delta}\tan^{d-1}(x)\Phi)_{,x}}{\tan^{d-1}(x)} - \frac{e^{-\delta}\mu^2\phi}{\cos^2(x)}, \quad (1.3)$$

where Π is the canonical momentum and $\Phi = \phi_{,x}$ is an auxiliary variable. The Einstein equation reduces to constraints, which can be written as

$$\delta_{,x} = -\sin(x)\cos(x)(\Pi^2 + \Phi^2) \quad (1.4)$$

$$M_{,x} = (\tan(x))^{d-1} \left[A \frac{(\Pi^2 + \Phi^2)}{2} + \frac{\mu^2\phi^2}{2\cos^2(x)} \right], \quad (1.5)$$

$$A = 1 - 2 \frac{\sin^2(x)}{(d-1)\tan^d(x)} \frac{M}{\tan^d(x)}, \quad (1.6)$$

where the mass function M asymptotes to the conserved ADM mass at the boundary $x = \pi/2$. We will restrict to $d = 4$ spatial dimensions. Since results are robust against changes in the type of initial data [?], we can take the initial data to be a Gaussian of the areal radius in the canonical momentum and trivial in the field. Specifically,

$$\Pi(t=0, x) = \epsilon \exp\left(-\frac{\tan^2(x)}{\sigma^2}\right), \quad \phi(t=0, x) = 0. \quad (1.7)$$

The width σ and field mass μ constitute the parameter space for our stability phase diagram.

We solve the Klein-Gordon evolution equations (1.2,1.3) and Einstein constraint equations (1.4,1.5) numerically using methods similar to those of [?] on a spatial grid of $2^n + 1$ grid points; we discuss the convergence properties of our code in the appendix. We denote the approximate horizon position x_H and formation time t_H by the first point such that $A(x_H, t_H) \leq 2^{7-n}$. In detail, we evolve the system in time using a 4th-order Runge-Kutta stepper and initially use a 4th-order Runge-Kutta

spatial integrator at resolution $n = 14$. If necessary, we switch to a 5th-order Dormand-Prince spatial integrator and increase resolution near horizon formation. Due to time constraints, we do not increase the resolution beyond $n = 21$ for any particular calculation; if a higher resolution would be required to track horizon formation for a given amplitude, we exclude that amplitude.

To determine the stability class of initial data with a given width σ , we allow evolutions to run to a maximum time of $t_{lim} = 500$ in AdS units, so t_{lim} is a lower limit for t_H for amplitudes that do not form a horizon within that time. Normally, however, if the initial data appears unstable, we only evolve amplitudes with $t_H \lesssim 0.6t_{lim}$; this is partly to save computational resources and partly to distinguish stable evolutions from collapsing ones. For unstable or metastable initial data, we find the best fit of the form $t_H = a\epsilon^{-p} + b$ to evolutions with $t_H > t_{fit}$, where t_{fit} is a constant time chosen such that amplitudes with evolutions that last longer are usually roughly perturbative;² in practice, $t_{fit} = 60$ gives results close to the perturbative result $p = 2$ for evolutions expected to be unstable from the literature, but we will also consider $t_{fit} = 80, 100$ as described below. In other words, since a given amplitude ϵ may be in the perturbative scaling regime for one set of initial data but nonperturbative for another, we compare initial data at similar horizon formation times (addressing the onset of perturbative behavior). Choosing t_{fit} as above gives consistent values of the fit parameters for the three values of t_{fit} for the largest and smallest initial data widths, which are unstable.

1.4 Phase Diagram of Stability

Here we give our main result, the phase diagram of stability classes as a function of initial profile width and scalar mass, along with a more detailed discussion of the scaling of horizon formation time with amplitude for varying initial data.

The stability phase diagram for spherically symmetric scalar field collapse in AdS_5 , treating the

²The power law plus constant fits the leading and first subleading contribution to t_H in a perturbative expansion in ϵ , and we have found that the subleading term is typically not negligible in the computationally accessible regime.

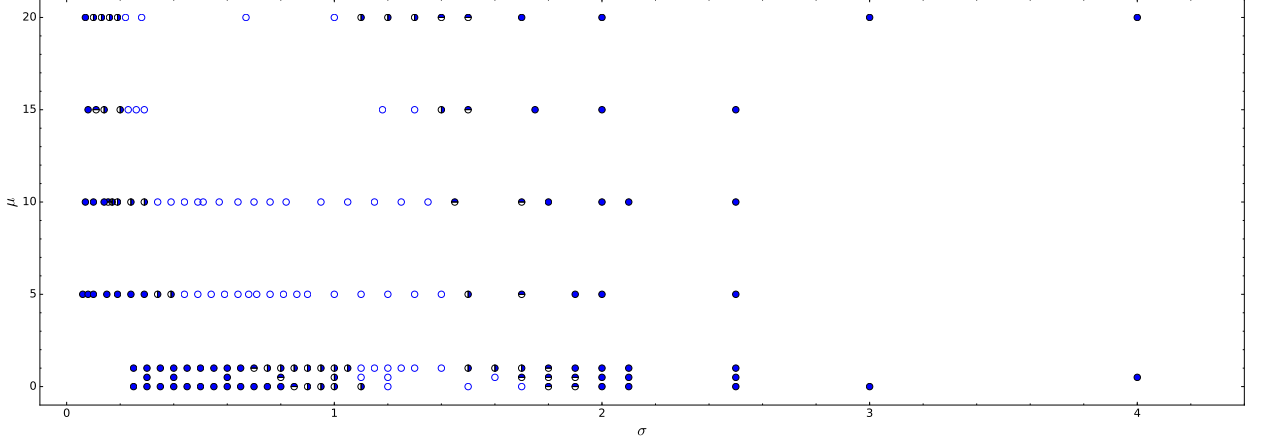


Figure 1.2: *Stability phase diagram as a function of initial data width σ and scalar mass μ . Filled circles represent unstable initial data, empty circles stable initial data, top-half-filled circles metastability, and right-half-filled circles irregular behavior.*

width σ of initial data and scalar field mass μ as tunable parameters, appears in figure 1.2. Each (μ, σ) combination that we evolved numerically is indicated by a circle, with filled and empty circles representing unstable and stable initial data respectively. The metastable class is represented by circles filled in the top half, while the irregular class are filled in the right half. At a glance, two features of the stability phase diagram are apparent: as μ increases, the island of stability moves toward smaller values of σ and takes up a gradually larger range of σ . To be specific, the stable class of initial data is centered at $\sigma = \bar{\sigma} \sim 1.4$ and has a width of $\Delta\sigma \sim 0.7$ for $\mu = 0, 0.5$, with $\bar{\sigma} \sim 1.2$ for $\mu = 1$. $\Delta\sigma$ increases to ~ 1.1 , and the island of stability is centered at $\bar{\sigma} \sim 0.9$ for $\mu = 5, 10$, while $\Delta\sigma \sim 1.2$ for $\mu = 15, 20$ with the stable class centered at $\bar{\sigma} \sim 0.8$. Note that the transition between “light field” and “heavy field” behavior occurs for $\mu > 1$ in AdS units.

The metastable and irregular classes appear at the shorelines of the island of stability, the boundary between unstable and stable classes. In particular, the slope of the power law $t_H \sim \epsilon^{-p}$ as $\epsilon \rightarrow 0$ increases as the width moves toward the island of stability, leading to metastable behavior. We find metastability at the large σ shoreline for all μ values considered and also at the small σ shoreline for several scalar masses. It seems likely that metastable behavior appears in only a narrow range of σ for larger μ , which makes it harder to detect in a numerical search, leading to its absence in

some parts of the stability phase diagram. We find irregular behavior at the small σ shoreline for every mass and at the large σ boundary for large μ , closer to stable values of σ than metastable initial data. This class of initial data includes a variety of irregular and non-monotonic behavior, as detailed below. Evidence for chaotic behavior especially becomes more prominent at larger values of μ , as we will discuss below.

1.4.1 Metastable versus unstable initial data

While stable and irregular initial data are typically apparent by eye in a plot of t_H vs ϵ , distinguishing the unstable from metastable classes is a quantitative task. As we described in section 1.3.2, we find the least squares fit of $t_H = a\epsilon^{-p} + b$ to all evolutions with $t_H > t_{fit}$ for the given (μ, σ) , running over values $t_{fit} = 60, 80, 100$. Using the covariance matrix of the fit, we also find the standard error for each fit parameter. We classify a width as having unstable evolution if the best fit value of p is within two standard errors of $p = 2$ for $t_{fit} = 60, 80$ or one standard error for $t_{fit} = 100$ (due to a smaller number of data points, the standard errors for $t_{fit} = 100$ tend to be considerably larger). In contrast, we classify a width as having metastable evolution if the best fit p is statistically significantly different from 2 (in that the best fit value is more than 2 standard errors from $p = 2$ for $t_{fit} = 60, 80$ and more than 1 standard error from $p = 2$ for $t_{fit} = 100$). This indicates that either further subleading contributions in a perturbative expansion of t_H are non-negligible in this regime for metastable initial data or that possibly metastable initial data are stable at the first nontrivial order in perturbation theory. Considering larger values of t_{fit} helps to ensure that the leading perturbative terms do not come to dominate for particular initial profile at the smallest computationally accessible amplitude values. In the case that the fit to $t_H = a\epsilon^{-p} + b$ has large reduced χ^2 or is sensitive to fitting algorithm, the data is not well-described by our fitting function, so we classify it as irregular (see the next subsection).

The fits $t_H = a\epsilon^{-p} + b$ allow us to explore the time scale of horizon formation across the stability phase diagram, for example through a contour plot of one of the coefficients vs σ and μ . In most cases, this has not been informative, but an intriguing feature emerges if we plot the normalization

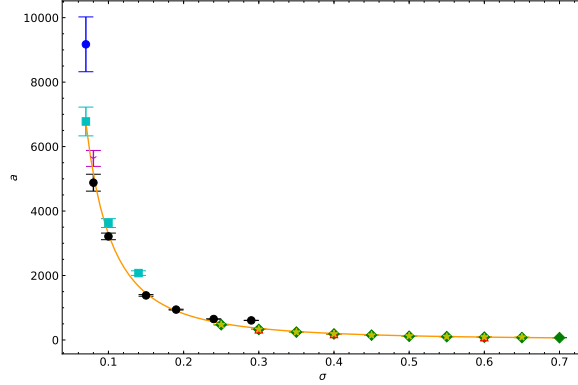
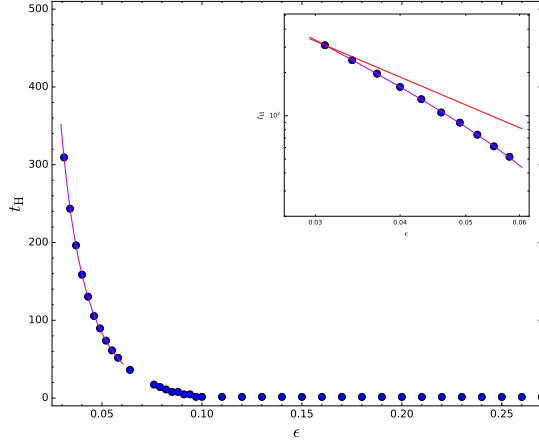


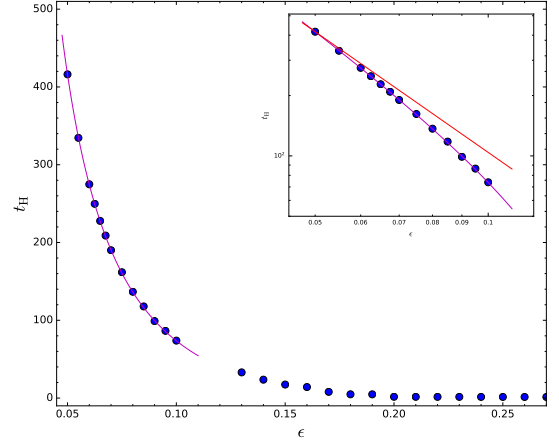
Figure 1.3: Coefficient a from the fit $t_H = a\epsilon^{-p} + b$ as a function of width σ using $t_{fit} = 60$. Shows data for $\mu = 0$ (green diamonds), 0.5 (red triangles), 1 (yellow stars), 5 (black circles), 10 (cyan squares), 15 (magenta Y), and 20 (blue circles). The orange line is the best power law fit.

coefficient a vs σ for unstable initial data at small σ , as shown in figure 1.3 for $t_{fit} = 60$. By eye, the coefficient is reasonably well described by the fit $a = (32.0 \pm 0.3)\sigma^{-(2.01 \pm 0.02)}$ (values following \pm are standard errors of the best fit values) *independent of scalar field mass*. This is not born out very well quantitatively; the reduced χ^2 for the fit is $\chi^2/\text{d.o.f.} = 180$, indicating a poor fit. However, the large χ^2 seems largely driven by a few outlier points with large scalar mass, so it is tempting to speculate that the gravitational collapse in this region of parameter space is driven by gradient energy, making all fields effectively massless at narrow enough initial σ . The picture is qualitatively similar if we consider the parameter a for $t_{fit} = 80, 100$ instead.

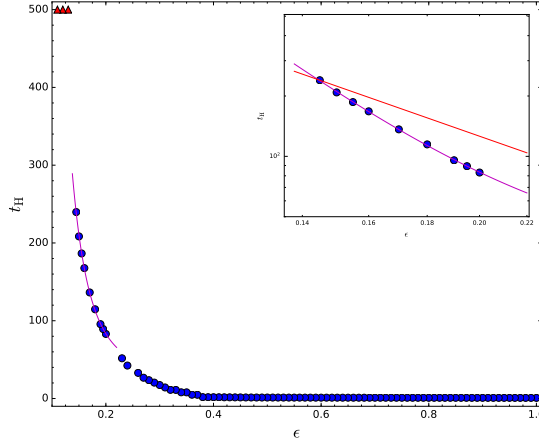
Several examples of metastable behavior appear in figure 1.4. These figures show both data from the numerical evolutions (blue dots and red triangles) and fits of the form $t_H = a\epsilon^{-p} + b$ for points with $t_H > t_{fit} = 60$ (magenta curves). The best fit parameters are given in table 1.1 along with the standard errors (listed following \pm for the fit values) and χ^2 values. The insets show the fit region with a log-log scale and an additional line (red) showing an ϵ^{-2} power law normalized to fit the smallest amplitude shown in the inset. It is visually clear that t_H grows faster than ϵ^{-2} for all these examples as ϵ decreases in the fit region (there is a significant constant offset in figure 1.4d).



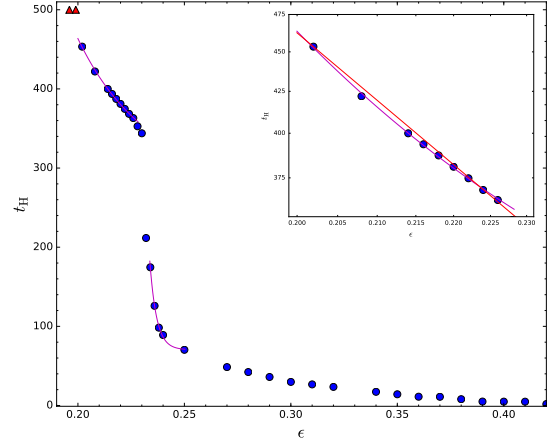
(a) $\mu = 15, \sigma = 1.5$



(b) $\mu = 5, \sigma = 1.7$



(c) $\mu = 0, \sigma = 1.8$



(d) $\mu = 0.5, \sigma = 1.7$

Figure 1.4: Metastable behavior: blue dots represent horizon formation and red triangles a lower limit on t_H . Magenta curves are fits $t_H = a\epsilon^{-p} + b$ over the shown range of amplitudes. Insets show the fit region with log-log scale; note that the fit is not strictly a power law, so the fits are not straight lines. See table 1.1 for best fit parameters. Red lines in insets are ϵ^{-2} power laws normalized to the t_H of the smallest amplitude shown.

Figures 1.4a, 1.4b demonstrate behavior typical of most of the instances of metastable initial data we have found; specifically, the initial data continue to collapse through horizon formation times of

	a	p	b	$\chi^2/\text{d.o.f.}$
$\mu = 15, \sigma = 1.5$	0.10 ± 0.01	2.33 ± 0.05	-27 ± 4	0.7736
$\mu = 5, \sigma = 1.7$	0.91 ± 0.06	2.07 ± 0.02	-33 ± 2	0.5070
$\mu = 0, \sigma = 1.8$	0.06 ± 0.02	4.3 ± 0.2	30 ± 5	1.502
$\mu = 0.5, \sigma = 1.7$ ($t_H < 0.4t_{lim}$)	$(4 \pm 32) \times 10^{-45}$	73 ± 5	70 ± 2	5.409
$(t_H > 0.72t_{lim})$	0.02 ± 0.03	5.6 ± 0.8	260 ± 20	1.078

Table 1.1: Best fit parameters for the cases shown in figure 1.4 restricting to $t_H > t_{fit} = 60$ and as noted. Listed errors (\pm values) are standard errors. $\chi^2/\text{d.o.f.}$ is the reduced χ^2 value used as a measure of goodness-of-fit.

$t_H \sim 0.6t_{lim}$ but with p significantly greater than the perturbative value of $p = 2$. Note that the evolutions of figure 1.4b have been extended to larger values of t_H to demonstrate that the evolutions continue to collapse to somewhat smaller amplitude values. Figure 1.4b is also of interest because its best fit value $p \approx 2.07 \pm 0.02$ is approximately as close to the perturbative value as several stable sets of initial data but has a smaller standard error for the fit, so the difference from the perturbative value is more significant (again, the value following the \pm is the standard error).

Figure 1.4c shows metastable evolution to $t_H \lesssim 0.6t_{lim}$ but then a sudden jump to stability until $t = t_{lim}$. In the figure, the fit has been extended to the largest non-collapsing amplitude, which demonstrates that there is no collapse over a time period significantly longer than the fit predicts. This example argues that metastable data may in fact become stable at the smallest amplitudes. On the other hand, figure 1.4d shows a similar jump in t_H to values $t_H < t_{lim}$; evolution at lower amplitudes shows metastable scaling with $p \approx 5.6 \pm 0.8$ for $360 < t_H < t_{lim}$. The figure also shows a metastable fit with larger reduced χ^2 at larger amplitudes corresponding to $t_{fit} < t_H < 0.4t_{lim}$. So this is another option: metastable behavior may transition abruptly to metastable behavior with different scaling (or possibly even perturbatively unstable behavior) at sufficiently small amplitudes. It is also reasonable to classify this case as irregular due to the sudden jump in t_H ; we choose metastable due to the clean metastable behavior at low amplitudes.

Our point of view is that initial data in the metastable class is distinct from the unstable class at finite amplitudes corresponding to $t_{fit} < t_H < 300$; they take longer to collapse at a fixed small value of ϵ than would be expected by the perturbative scaling. An alternate point of view is to ask whether we can determine if a given set of initial data is perturbatively unstable in the $\epsilon \rightarrow 0$ limit. We have already seen that metastable initial data does not follow the perturbative scaling when fit to $t_H = a\epsilon^{-p} + b$, the first two terms of the perturbative expansion. However, it is possible that a perturbative description applies but requires a further subleading term. To test this hypothesis, we fit unstable and metastable initial data to $t_H = a\epsilon^{-p} + b + c\epsilon^2$; as described earlier in this section, we determine if p is within two standard errors of the perturbative value $p = 2$ (or one standard error for $t_{fit} = 100$).

The unstable class of initial data is instructive. For the new fits of unstable initial data, p is statistically equal to 2, and the new values of a, p, b are consistent with the values from the old fits to within two standard errors (or sometimes slightly more). The fit value of c is uniformly within a standard error of zero, and, for the amplitude values in the fit region, the ϵ^2 term is small compared to the constant and ϵ^{-2} terms. What is more, for some unstable initial data near the island of stability, the original $t_H = a\epsilon^{-p} + b$ fits for $t_{fit} = 60$ have $p > 2$ statistically; on the other hand, the new fits have $p = 2$ within statistical error. In other words, the perturbative expansion is still valid but requires more terms. Part of the metastable class of initial data also behaves in this manner and could therefore be reasonably considered to be perturbatively unstable. Of the metastable initial data we found, these are $\sigma = 1.9$ for $\mu = 0$, $\sigma = 0.8$ and 1.9 for $\mu = 0.5$, $\sigma = 0.7$ for $\mu = 1$, $\sigma = 1.7$ for $\mu = 5$, $\sigma = 0.155$ for $\mu = 10$, $\sigma = 0.11$ and 1.5 for $\mu = 15$, and $\sigma = 1.5$ for $\mu = 20$. In addition, $\mu = 1, \sigma = 1.8$ and $\mu = 10, \sigma = 1.7$ initial data have similar behavior, but p is not statistically consistent with 2 for any of the fit regions, though it is closer than in the original fits. On the other hand, the other metastable initial data ($\sigma = 0.85$ and 1.8 for $\mu = 0$, $\sigma = 1.7$ and 1.8 for $\mu = 0.5$, $\sigma = 1.45$ for $\mu = 10$, and $\sigma = 1.4$ for $\mu = 20$) show no evidence for perturbative behavior. Specifically, p remains statistically larger than 2 for all fits, the ϵ^2 term in the new fit is roughly the same magnitude as the other terms, and the a, p, b values in the new fits are not

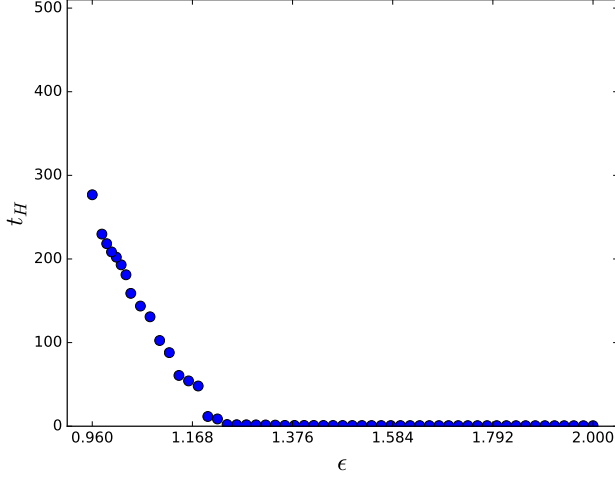
statistically consistent with the original fits.

1.4.2 Irregular behaviors

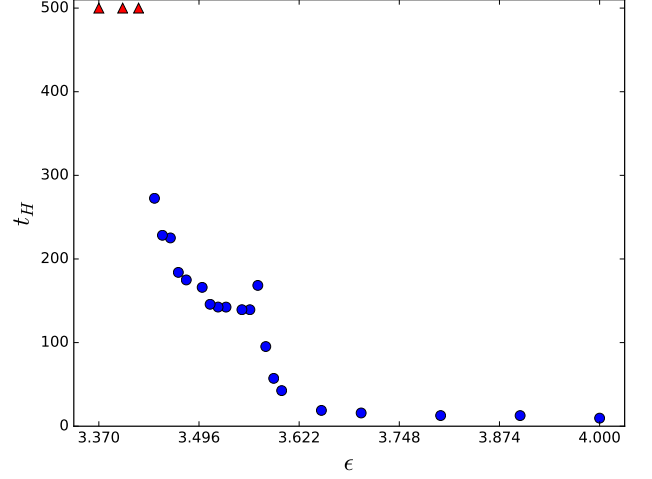
We have found a variety of irregular behaviors at the transition between the metastable and stable classes which we have classified together as irregular initial data; however, it may be better to describe them as separate classes. The stability phase diagram 1.2 indicates that the irregular class extends along the “inland” side of the small σ shoreline and at least part of the large σ shoreline of the island of stability. What is not clear from our evolutions up to now is whether each type of behavior appears along the entire shoreline or if they appear in pockets at different scalar field masses. Examples of each type of behavior that we have found appear in figure 1.5.

The first type of irregular behavior, shown in figure 1.5a, is monotonic (t_H increases with decreasing ϵ as usual), but it is not well fit by a power law. In fact, this behavior would classify as metastable by the criterion of section 1.4.1 in that the power law of the best fit $t_H = a\epsilon^{-p} + b$ is significantly different from $p = 2$, except for the fact that the reduced χ^2 value for the fit is very large (greater than 10) and also that different fitting algorithms can return significantly different fits, even though the data may appear to the eye like a smooth power law. In any case, this type of behavior apparently indicates a breakdown of metastable behavior and hints at the appearance of non-monotonicity. So far, our evolutions have not demonstrated sudden jumps in t_H typical of stability at low amplitudes, however.

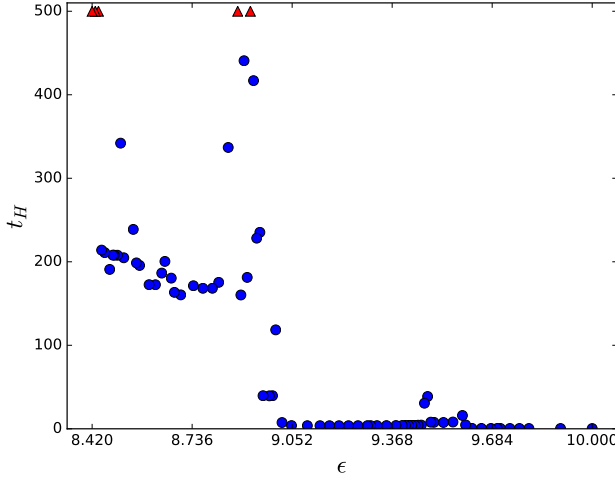
Figure 1.5b exemplifies non-monotonic behavior in the irregular class. This type of behavior, which was noted already by [?], involves one or more sudden jumps in t_H as ϵ decreases, which may be followed by a sudden decrease in t_H and then resumed smooth monotonic increase in t_H . There are suggestions that this type of initial data is stable at low amplitudes due to the usual appearance of non-collapsing evolutions, but it is worth noting that these amplitudes could instead experience another jump and decrease in t_H , just at $t_H > t_{lim}$. Finally, [?] studied this type of behavior in some detail, denoting it as “quasi-stable.”



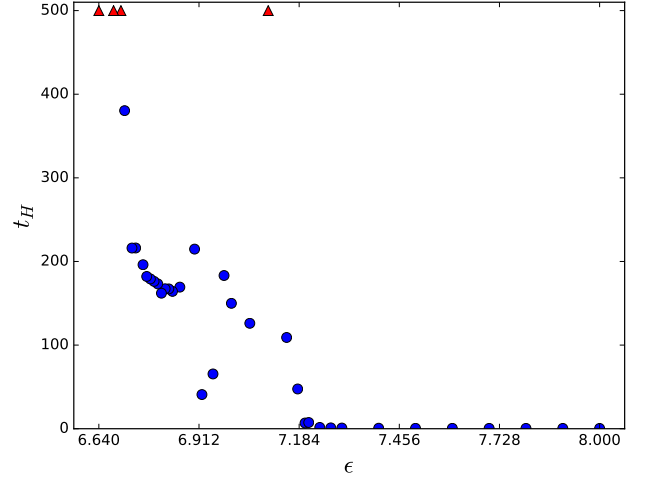
(a) $\mu = 0.5, \sigma = 1$



(b) $\mu = 5, \sigma = 0.34$



(c) $\mu = 20, \sigma = 0.16$



(d) $\mu = 20, \sigma = 0.19$

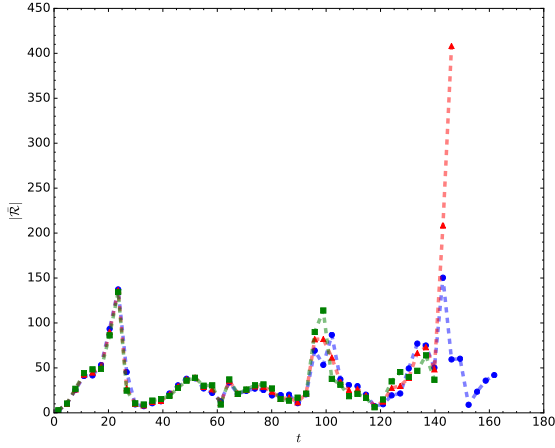
Figure 1.5: *Irregular behavior: blue dots represent horizon formation and red triangles a lower limit on t_H .*

Some irregular initial data demonstrates evidence of chaotic behavior, in that t_H appears to be sensitive to initial conditions (ie, value of amplitude) over some range of amplitudes. This type of behavior appears over the range of masses (see figure 1.1d for a mild case for massless scalars), but it is more common and more dramatic at larger μ . Figures 1.5c, 1.5d represent the most extreme behavior of this type among the initial data that we studied with collapse at $t_H < 50$ not very far separated from amplitudes that do not collapse for $t < t_{lim}$ along with an unpredictable pattern of

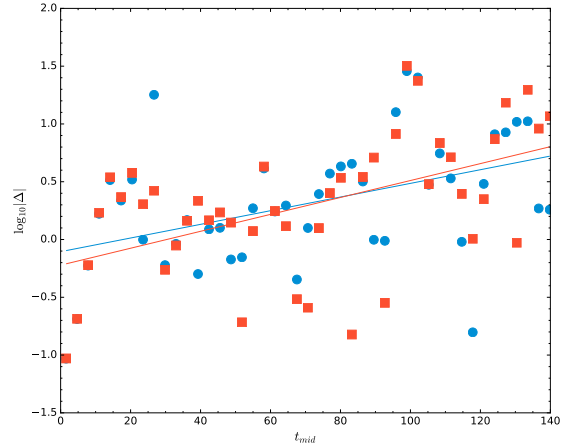
variation in t_H . This type of evidence for chaotic behavior has been seen previously in the collapse of transparent but gravitationally interacting thin shells in AdS [?] as well as in the collapse of massless scalars in AdS₅ Einstein-Gauss-Bonnet gravity [?,?]; these references speculated that the t_H vs ϵ curve is fractal. In both cases, this type of behavior is hypothesized to be due to the transfer of energy between two infalling shells, with horizon formation only proceeding when one shell is sufficiently energetic. In the latter case, the extra scale of the theory (given by the coefficient of the Gauss-Bonnet term in the action) leads the single initial pulse of scalar matter to break into two pulses.

We should therefore ask two questions: does this irregular behavior show evidence of true chaos, and is a similar mechanism at work here? We note first that [?] found evidence (using a modified box test) that the t_H vs ϵ curve has a non-integer fractal dimension for plots visually similar to our figures 1.5c, 1.5d. Here, to quantify the presence of chaos, we examine the difference in time evolution between similar initial conditions (nearby amplitudes), which diverge exponentially in chaotic systems. Specifically, any quantity Δ should satisfy $|\Delta| \propto \exp(\lambda t)$ for Lyapunov coefficient λ . Our characteristic will be the upper envelope of the Ricci scalar at the origin per light crossing time, $\bar{\mathcal{R}}(t)$. We consider three sets of irregular initial data: a massless scalar of width $\sigma = 1.1$ with amplitudes $\epsilon = 1.02, 1.01, 1.00$ (see figure 1.1d), a $\mu = 5$ massive scalar of width $\sigma = 0.34$ and $\epsilon = 3.52, 3.51, 3.50$, and a $\mu = 20$ scalar of width $\sigma = 0.19$ and $\epsilon = 6.98, 6.95, 6.92$ (figure 1.5d). We also calculated determined the Lyapunov coefficient for unstable initial data with $\mu = 0.5, \sigma = 0.3$, and $\epsilon = 1.22, 1.20, 1.18$ for comparison.

Figure 1.6 details evidence for chaotic evolution in the $\mu = 5, \sigma = 0.34$ case; figure 1.6a shows our characteristic function $\bar{\mathcal{R}}(t)$ for the amplitudes $\epsilon_1 = 3.50, \epsilon_2 = 3.51$, and $\epsilon_3 = 3.52$. By eye, $\bar{\mathcal{R}}$ shows noticeable differences after a long period of evolution. These are more apparent in figure 1.6b, which shows the log of the differences $\Delta_{ab} \equiv \bar{\mathcal{R}}_{\epsilon_a} - \bar{\mathcal{R}}_{\epsilon_b}$, along with the best fits. Although there is considerable noise — or oscillation around exponential growth — in the differences (leading to R^2 values $\sim 0.2, 0.26$ for the fits), the average slope gives Lyapunov coefficient $\lambda = 0.007$ (within the error bar of each slope), and each slope differs from zero by more than 3 standard errors. One



(a) Upper envelope of Ricci scalar at origin



(b) $\log |\Delta|$ vs. t_{mid}

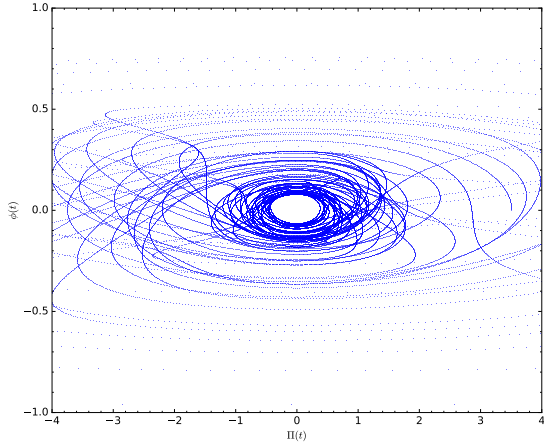
Figure 1.6: Left: The upper envelope of the Ricci scalar for amplitudes $\epsilon_1 = 3.50$ (blue circles), $\epsilon_2 = 3.51$ (red triangles), and $\epsilon_3 = 3.52$ (green squares) for $\mu = 5, \sigma = 0.34$. Right: $\log(|\Delta_{12}|)$ and best fit (blue circles and line) and $\log(|\Delta_{23}|)$ and best fit (red squares and line), calculated as a function of the midpoint t_{mid} of the time interval.

interesting point is that the t_H vs ϵ curve in figure 1.5b does not appear chaotic to the eye, even though it shows some of the mathematical signatures of chaos at least for $\epsilon_1 < \epsilon < \epsilon_3$ (the visible spike in t_H is at $\epsilon \sim 3.57$).

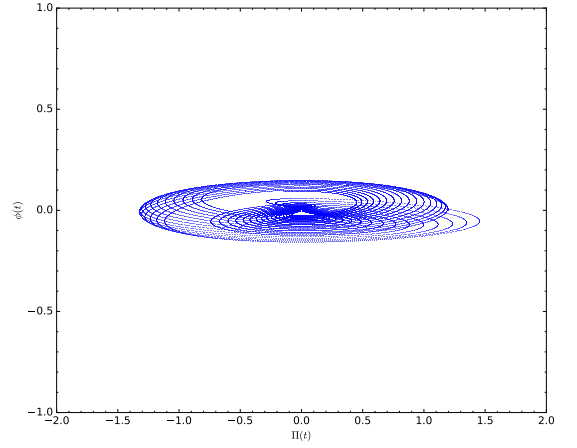
The story is similar for the massless and $\mu = 20$ cases we studied, which exhibit λ values that differ from zero by at least 1.9 standard errors; see table 1.2. This is a milder version of the behavior noted by [?, ?, ?], especially for the $\mu = 5$ case studied. One thing to note is that the strength of oscillation in $\log(|\Delta|)$ around the linear fit increases with increasing mass, so that the two best fit Lyapunov exponents for $\mu = 20$ are no longer consistent with each other at the 1-standard deviation level. We should note, however that the unstable initial data with $\mu = 0.5, \sigma = 0.3$ also exhibits a statistically positive Lyapunov exponent, though we should note that the value of λ quoted in table 1.2 includes the time shortly before horizon formation, which does increase λ somewhat (though not more than the quoted error).

		λ	average λ
$\mu = 0, \sigma = 1.1$	Δ_{12}	0.011 ± 0.005	0.011
	Δ_{23}	0.011 ± 0.005	
$\mu = 0.5, \sigma = 0.3$	Δ_{12}	0.021 ± 0.0007	0.022
	Δ_{23}	0.024 ± 0.001	
$\mu = 5, \sigma = 0.34$	Δ_{12}	0.006 ± 0.002	0.007
	Δ_{23}	0.007 ± 0.002	
$\mu = 20, \sigma = 0.19$	Δ_{12}	0.046 ± 0.009	0.032
	Δ_{23}	0.019 ± 0.007	

Table 1.2: Best fit Lyapunov coefficients λ for adjacent amplitude pairs and average λ value for each μ, σ system studied. Standard errors are given following \pm signs.



(a) $\mu = 5, \sigma = 0.34, \epsilon = 3.51$



(b) $\mu = 0.5, \sigma = 0.3, \epsilon = 1.20$

Figure 1.7: Trajectories in $\Pi(x=0), \phi(x=0)$ phase space for one irregular and one unstable evolution. Trajectories are shown for $t < 50$.

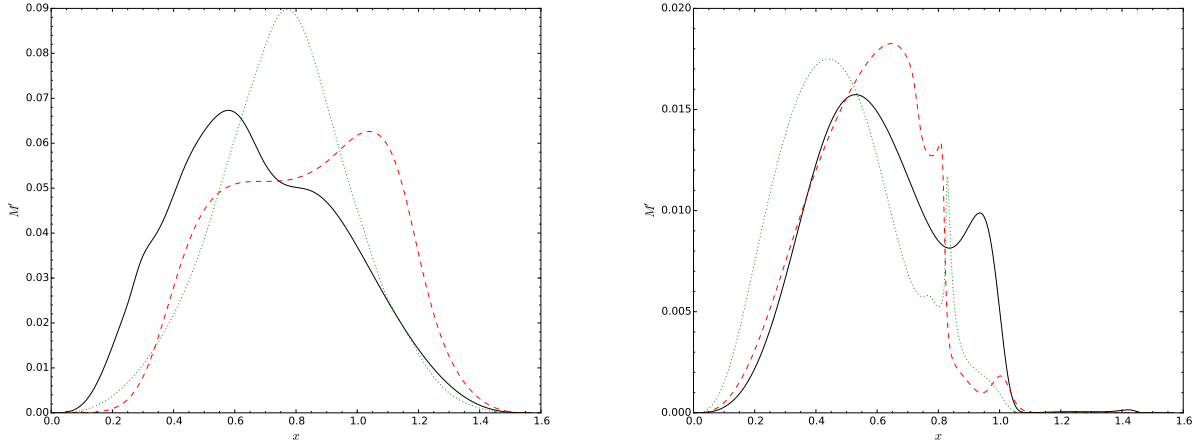
Since the Lyapunov coefficients do not distinguish the irregular and unstable cases, we also consider the phase space trajectories of the evolutions. Following [?], we consider the trajectory of evolutions in Π and ϕ evaluated at the origin for $t \leq 50$ in figure 1.7. Neither the $\mu = 5, \sigma = 0.34, \epsilon = 3.51$ (figure 1.7a) or $\mu = 0.5, \sigma = 0.3, \epsilon = 1.20$ (figure 1.7b) trajectories close, though there is a clear

difference. Specifically, the former trajectory is visually disorganized (that is, strongly varying orbits) with very rapid motion (seen in the gap between points on the trajectory between plotted time steps). Meanwhile, the latter motion is comparatively regular, typical of quasi-periodic motion. Figure 1.7a is typical of turbulence and clearly shows that these evolutions are nonperturbative, even though t_H is large (well into the perturbative regime for unstable initial data).

To sum up, we have identified irregular initial data that shows evidence of chaotic behavior. Specifically, several of the t_H vs ϵ curves appear qualitatively similar to analogous plots in [?, ?, ?], which were demonstrated to have fractal-like behavior (including fractional fractal dimension in one case). Furthermore, a number of cases of irregular initial data (and some unstable) have positive Lyapunov exponents; phase space trajectories for irregular initial data show very rapid motion typical of turbulence, while unstable initial data have more regular trajectories. Taken together, this is strong evidence for chaotic behavior for some irregular initial data, similar to that discussed in other studies of gravitational collapse in AdS. Furthermore, this is the first evidence of chaos in the t_H vs ϵ curve for gravitational collapse of a massless scalar in AdS to our knowledge.

The mechanism underlying the possibly chaotic behavior seems somewhat different or at least weaker than the two-shell or Einstein-Gauss-Bonnet systems. When examining the time evolution of the mass distributions of these data, we see a single large pulse of mass energy that oscillates between the origin and boundary without developing a pronounced peak. However, there is also apparently a smaller wave that travels across the large pulse. We can see this by comparing snapshots of the mass distribution at different times, as in figure 1.8. In the massless case examined, this wave deforms the pulse, leading to a double-shoulder appearance seen at two times in figure 1.8a. In the $\mu = 5, \sigma = 0.34$ case, the secondary wave is more like a ripple, usually smaller in amplitude but more sharply localized, as toward the right side of the main pulse in figure 1.8b. So the chaotic behavior may be caused by the relative motion of the two waves, rather than energy transfer between two shells. In this hypothesis, a horizon would form when both waves reach the neighborhood of the origin at the same time.

As a note, we have run convergence tests on several sets of irregular initial data and find that



(a) $\mu = 0, \sigma = 1.1, \epsilon = 1.01$, at times $t = 60$ (solid black), $t = 62$ (dashed red), $t = 64$ (dotted green) (b) $\mu = 5, \sigma = 0.34, \epsilon = 3.52$, at times $t = 132$ (solid black), $t = 137$ (dashed red), $t = 140$ (dotted green)

Figure 1.8: Radial derivative of the mass function at the indicated time for two systems that show evidence of chaos. Note the appearance of a secondary wave on top of the main pulse. (μ, σ, ϵ) as indicated.

our calculations are convergent overall, as expected (even at lower resolution than we used). In particular, the massless scalar evolutions studied in table 1.2 are convergent already at resolution given by $n = 12$ (note that we typically start at $n = 14$); we also observe convergent behavior for the $\mu = 5$ evolutions discussed in table 1.2. We have therefore validated that nonmonotonic behavior and even evidence of chaos occurs. The only caveat may be for some of the apparently initial data with scalar mass $\mu = 20$, which nonetheless appear well-behaved according to other indicators. The reader may or may not wish to take them at face value but should recall that we have presented other chaotic initial data with rigorously convergent evolutions. See the appendix for a more detailed discussion.

1.5 Spectral analysis

As we discussed in the introduction, instability toward horizon formation proceeds through a turbulent cascade of energy to shorter wavelengths or, more quantitatively, to 1st-order scalar eigenmodes with more nodes. Inverse cascades are typical of stable evolutions. Therefore, understanding the energy spectrum of our evolutions, both initially and over time, sheds light on the behavior of the self-gravitating scalar field in asymptotically AdS spacetime, providing a heuristic analytic understanding of the stability phase diagram.

The (normalizable) eigenmodes e_j are given by Jacobi polynomials as

$$e_j(x) = \kappa_j \cos^{\lambda_+}(x) P_j^{(d/2-1, \sqrt{d^2+4\mu^2}/2)}(\cos(2x)) \quad (1.8)$$

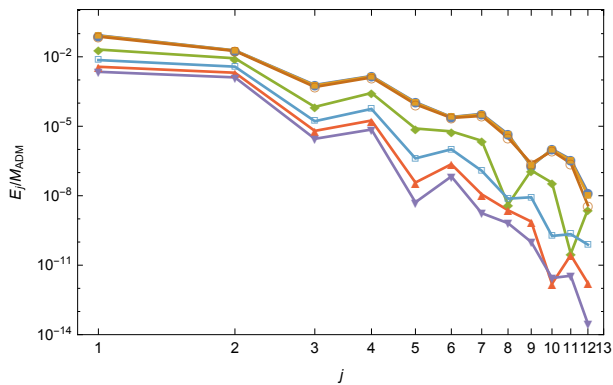
(κ_j is a normalization constant) with eigenfrequency $\omega_j = 2j + \lambda_+$ and $\lambda_+ = (d + \sqrt{d^2 + 4\mu^2})/2$ in AdS_{d+1} for $j = 0, 1, \dots$ (see [?, ?] for reviews). Including gravitational backreaction, we define the energy spectrum

$$E_j \equiv \frac{1}{2} \left(\Pi_j^2 - \phi_j \ddot{\phi}_j \right), \quad (1.9)$$

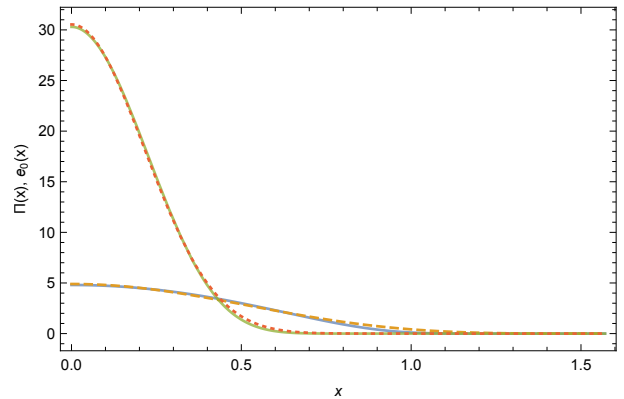
where

$$\begin{aligned} \Pi_j &= \left(\sqrt{A} \Pi, e_j \right), \quad \phi_j = (\phi, e_j), \\ \ddot{\phi}_j &= (\cot^{d-1}(x) \partial_x [\tan^{d-1}(x) A \Phi] - \mu^2 \sec^2(x) \phi, e_j), \end{aligned} \quad (1.10)$$

and the inner product is $(f, g) = \int_0^{\pi/2} dx \tan^{d-1}(x) f g$. The sum of E_j over all modes is the conserved ADM mass.



(a) Best fit gaussian energy spectra.



(b) Best fit gaussian and zeroth eigenmode.

Figure 1.9: Left: Spectra of the best fit gaussians (1.7) to the $j = 0$ eigenmode for masses $\mu = 0$ (blue circles), 0.5 (yellow squares), 1 (empty orange circles), 5 (green diamonds), 10 (empty cyan squares), 15 (upward red triangles), and 20 (downward purple triangles). Right: an overlay of the best fit Gaussian and e_0 eigenmode for $\mu = 0$ (solid blue is best fit, orange dashed is eigenmode) and $\mu = 20$ (solid green, red short dashes).

1.5.1 Dependence on mass

The most visibly apparent feature of the stability phase diagram of figure 1.2 is that the island of stability both expands and shifts to smaller widths as the scalar mass increases. As it turns out, the energy spectrum of the Gaussian initial data (1.7) provides a simple heuristic explanation.

It is well established both in perturbation theory and numerical studies that initial data given by a single scalar linear-order eigenmode is in fact nonlinearly stable, and the spectra of many quasi-periodic solutions are also dominated by a single eigenmode. As a result, we should expect Gaussian initial data that approximates a single eigenmode (which must be $j = 0$ due to lack of nodes) to be stable. To explore how this depends on mass, we find the best fit values of ϵ, σ for the $j = 0$ eigenmode for each mass that we consider (defined by the least-square error from the Gaussian to a discretized eigenmode); this is the “best approximation” Gaussian to the eigenmode. Then we find the energy spectrum of that best-fit Gaussian; these are shown in figure 1.9a. From the figure, it is

clear that the $j = 0$ eigenmode is closer to a Gaussian at larger masses. That is, other eigenmodes contribute less to the Gaussian's spectrum at higher masses (by several orders of magnitude over the range from $\mu = 0$ to 20). Simply put, the shape of the $j = 0$ eigenmode is closer to Gaussian at higher masses, which suggests that the island of stability should be larger at larger scalar field mass. Figure 1.9b compares the $j = 0$ eigenmode and best fit Gaussian for $\mu = 0$ and 20; on inspection, there is more deviation between the eigenmode and Gaussian for the massless scalar.

In addition, the best-fit Gaussian width decreases from $\sigma \sim 0.8$ for a massless scalar as the mass increases. At $\mu = 20$, the best-fit width is $\sigma \sim 0.31$. This suggests that Gaussians that approximate the $j = 0$ mode well enough are narrower in width at higher masses. An interesting point to note is that the island of stability for $\mu = 0, 0.5$ is actually centered at considerably larger widths than the best-fit Gaussian. This may not be surprising, since the best-fit Gaussians at low masses actually receive non-negligible contributions from higher mode numbers; moving away from the best-fit Gaussian can actually reduce the power in higher modes. For example, the stable initial data shown in figure 1.10a below has considerably less power in the $j = 2$ mode.

1.5.2 Spectra of different behaviors

A key question that one might hope to answer is whether the stability class of a given (μ, σ) can be determined easily by direct inspection of the initial data without requiring many evolutions at varying amplitudes. The initial energy spectra for examples of each class, including monotonic, non-monotonic, and apparently chaotic irregular behaviors, are shown in figure 1.10. These spectra are taken from among the smallest amplitudes we evolved in order to minimize backreaction effects.

Unfortunately, the initial energy spectra do not seem to provide such a method for determining the stability class. Very broadly speaking, stable and metastable (μ, σ) correspond to initial spectra that drop off fairly quickly from the $j = 0$ mode as j increases, while unstable and irregular behaviors tend to have roughly constant or even slightly increasing spectra up to $j = 5$ or 10. However, figure 1.10d shows that some irregular initial data have spectra that decrease rapidly after a small increase

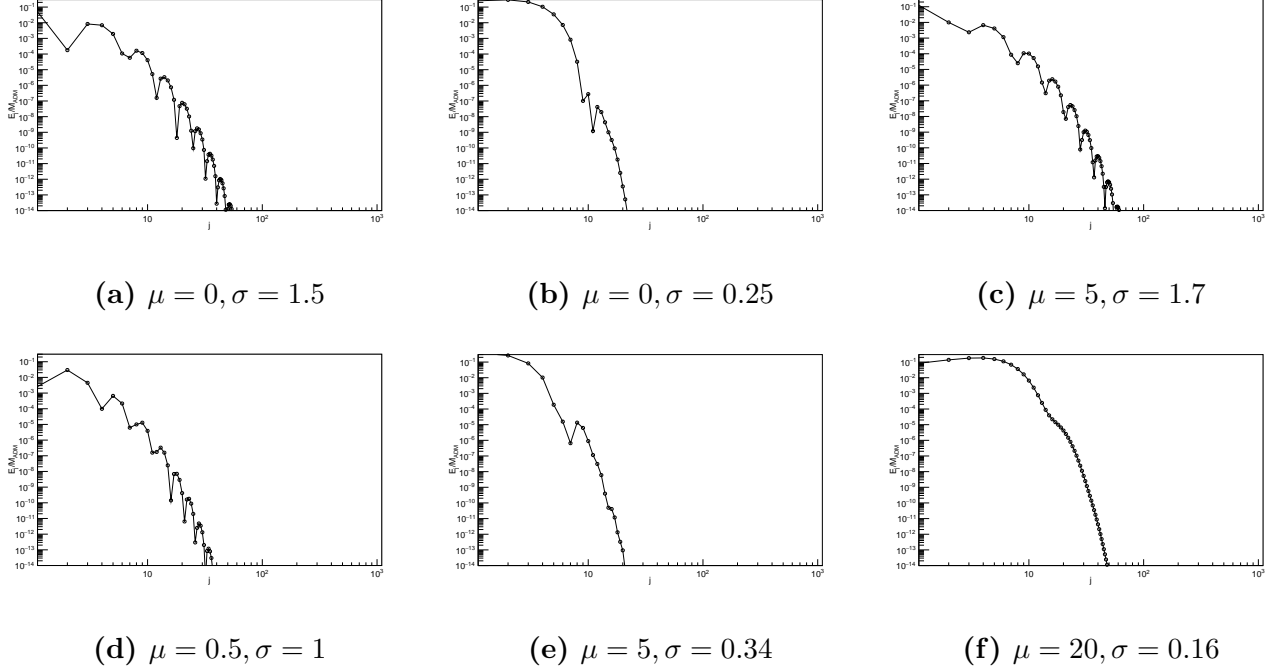


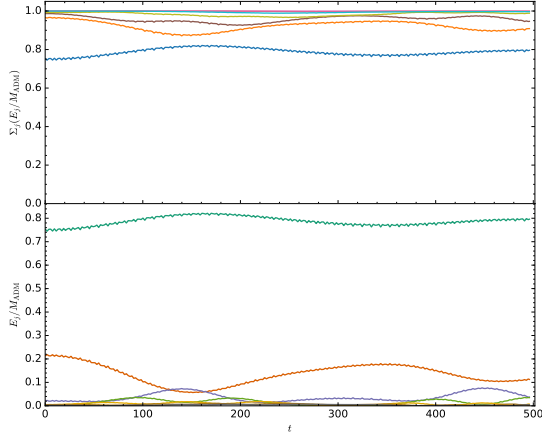
Figure 1.10: Initial ($t = 0$) energy spectra for the indicated evolutions. In order, these represent stable, unstable, metastable, monotonic irregular, non-monotonic irregular, and chaotic irregular initial data.

from $j = 1$ to $j = 2$. Kinks in the spectrum are more prevalent for widths of the AdS scale or larger, while spectra for smaller widths tend to be smoother.

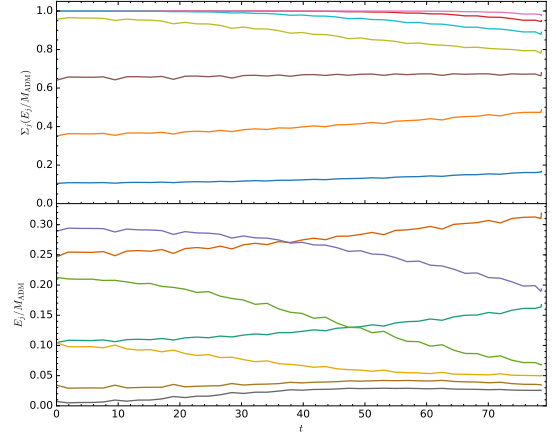
1.5.3 Evolution of spectra

While the initial spectrum for a given (μ, σ) pair does not have predictive value regarding the future behavior as far as we can tell, the time dependence of the spectrum throughout the evolution of the system is informative. Figure 1.11 shows the time-dependence of spectra for examples of the stable, unstable, metastable, and chaotic irregular classes. In each figure, the lower panel shows the fraction E_j/M_{ADM} in each mode up to $j = 6$, while the upper panel shows the cumulative fraction $\sum_j E_j/M_{ADM}$ to the mode 2^k with $k = 0$ to 5.

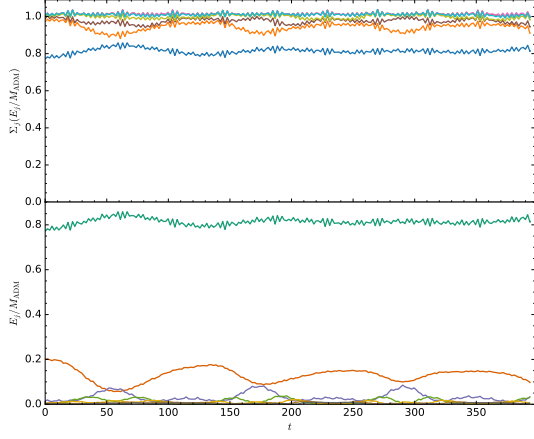
The difference between stable evolution in figure 1.11a and unstable evolution in figure 1.11b is readily apparent. As the evolution proceeds, we expect a cascade of energy into higher mode



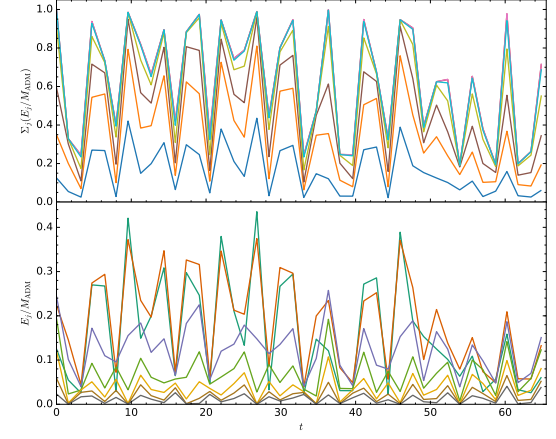
(a) $\mu = 0, \sigma = 1.8, \epsilon = 0.13$



(b) $\mu = 0, \sigma = 0.25, \epsilon = 2.28$



(c) $\mu = 0.5, \sigma = 1.7, \epsilon = 0.216$



(d) $\mu = 20, \sigma = 0.19, \epsilon = 6.95$

Figure 1.11: *The time dependence of the energy spectra as a fraction of the total ADM mass for the indicated μ, σ, ϵ . Lower panels show the lowest 7 modes (in colors cyan, red, purple, green, yellow, brown, and gray respectively). Upper panels show cumulative energy to mode $j = 0, 1, 2, 4, 8, 16, 32$ (in colors blue, orange, brown, yellow, aqua, red, and magenta).*

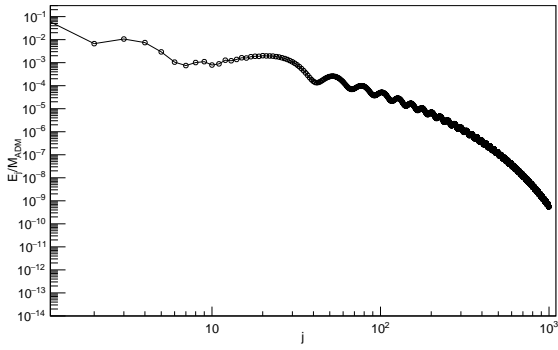
numbers, but inverse cascades to lower modes can also occur. The stable evolution shows a slow pattern of cascades and inverse cascades, in fact. On the other hand, the unstable evolution shows a nearly monotonic cascade of energy into the highest modes along with a simultaneous cascade

of energy into the lowest modes (therefore depleting intermediate modes). These are common observations in the literature and are included here for completeness.

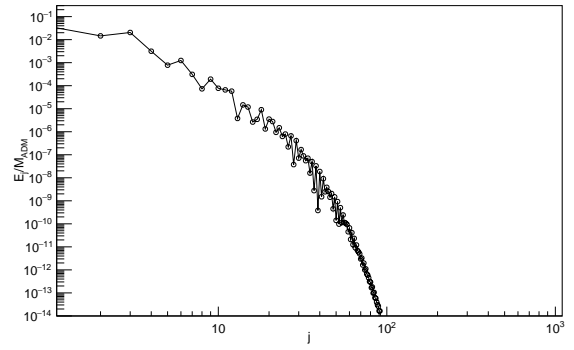
The metastable evolution shown in figure 1.11c is interesting in light of the stable and unstable spectra. The amplitude shown is from the “unstable” portion of figure 1.4d, the part consistent with the perturbative scaling $t_H \sim \epsilon^{-2}$. However, the spectrum shows a similar pattern of slow cascades and inverse cascades to the stable initial data example, though on a somewhat faster time scale in this case. While perhaps surprising, this is in keeping with the similarities noted between the initial spectra in figures 1.10a and 1.10c. We have also checked that the time-dependent spectrum at a higher amplitude with $t_H \sim 100$ follows the same pattern as 1.11c; in fact, it looks essentially the same but simply ends at an earlier time. This lends some credence to the idea that metastable initial data is stable at lowest nontrivial order in perturbation theory, with instability triggered by higher-order corrections. Alternately, the instability could be caused by an oscillatory singularity in the perturbative theory, as discussed in [?, ?, ?, ?] in the case of two-mode initial data. These divergences do not appear in the energy spectrum.

Figure 1.11d shows the time-dependence of the spectrum in an irregular evolution, specifically $\mu = 20, \sigma = 0.19$ at $\epsilon = 6.95$, which is in the chaotic region of the t_H vs ϵ plot in figure 1.5d. There is rapid energy transfer among modes, including cascades out of and inverse cascades into mode numbers $j \leq 32$ over approximately a light-crossing time. It is easy to imagine that horizon formation might occur at any of the cascades of energy into higher modes, leading to seemingly random jumps in t_H as a function of amplitude.

Finally, the time-evolved energy provide another possible measure of approximate thermalization in the dual CFT; namely, the spectrum should approach an (exponentially cut-off) power law at thermalization. In most cases, this occurs shortly before horizon formation, but there are exceptions, such as the late time behavior of initial data below the critical mass for black hole formation in Einstein-Gauss-Bonnet gravity [?]. When there is evidence of chaotic behavior, it is particularly interesting to know if the spectra for similar amplitudes approach a power law at similar times even if horizons form at very different times. Figure 1.12 shows the energy spectra for two amplitudes



(a) $\epsilon = 1.01$



(b) $\epsilon = 1.02$

Figure 1.12: Spectra at time $t \approx 71$ for $\mu = 0, \sigma = 1.1$ for the two amplitudes given. $\epsilon = 1.01$ forms a horizon at $t_H \approx 71.1$, $\epsilon = 1.02$ at $t_H \approx 248.0$.

in the chaotic region of the t_H vs ϵ plot for $\mu = 0, \sigma = 1.1$. Figure 1.12a is the spectrum just before horizon formation for $\epsilon = 1.01$, while figure 1.12b is the spectrum at approximately the same time for $\epsilon = 1.02$, which is very long before horizon formation. In this example, we see that the spectrum does approach a power law for the evolution that is forming a horizon, while the other evolution demonstrates a more rapid decay (typically fit by a power law times an exponential in the literature). Therefore, this example suggests that a power law spectrum may yield similar results to horizon formation as a measure of thermalization in the dual CFT.

1.6 Discussion

For the first time, we have presented the phase diagram of stability of AdS_5 against horizon formation, treating the scalar field mass μ and width σ of initial data as free parameters. In addition to mapping the location of the so-called “island of stability,” we have gathered evidence for two non-perturbative classes on the “shorelines” of the island, the metastable and irregular classes. While these must either exhibit stability (no collapse below some critical amplitude) or instability (collapse at arbitrarily small but finite amplitude) as the amplitude $\epsilon \rightarrow 0$, they are distinguished

by their behavior at computationally accessible (finite) amplitudes. While perturbatively unstable evolutions obey $t_H \propto \epsilon^{-2}$ as $\epsilon \rightarrow 0$ (and show evidence of this behavior at finite ϵ), metastable initial data follows $t_H \propto \epsilon^{-p}$ for $p > 2$ over a range of amplitudes $\epsilon > 0$. The irregular class is characterized by horizon formation times t_H that are not well described by a power law and sometimes exhibit non-monotonicity or even evidence of chaos. Both of these classes appear across the range of μ values that we study and at both small- and large-width boundaries of the stable class of initial data.

At this time, it is impossible to say whether metastable initial data is stable or unstable as $\epsilon \rightarrow 0$ (or if all metastable data behaves in the same way in that limit). Our numerical evolutions include cases in which the lowest amplitudes jump either to metastable scaling with smaller p or to evolutions that do not collapse over the timescales we study. In many cases, too, the power law $t_H \propto \epsilon^{-p}$ with p some fixed value > 2 is robust as we exclude larger amplitudes from our fit. It is also possible that some metastable initial data is stable in the perturbative theory (ie, to ϵ^3 order in a perturbative expansion) but not at higher orders.

The irregular class seems likely to be (mostly) stable at arbitrarily small amplitudes based on our numerical evolutions, though we have not found a critical amplitude for monotonic irregular initial data. The irregular initial data includes the “quasi-stable” initial data described in [?, ?], which has a sudden increase then decrease in t_H as ϵ decreases as well as evidence for chaotic behavior. In fact, we have found evidence for weakly chaotic behavior for non-monotonic initial data in the form of a small but nonzero Lyapunov coefficient and in the phase space trajectory. Both non-monotonicity and chaos become stronger and more common at larger scalar masses; however, we have also found evidence of chaotic behavior for the massless scalar including in the t_H vs ϵ curve. To our knowledge, this is the first evidence of chaos in this relationship for spherically symmetric massless scalar collapse in AdS, which is particularly interesting because there is only one physically meaningful ratio of scales, σ as measured in AdS units.

While we have emphasized the appearance of new behaviors outside perturbation theory, metastable and irregular initial data are interesting potential subjects for analysis in the multiscale perturbation

theory. A key question is if they demonstrate any unusual behavior there or map directly onto the stable or unstable classes.

Aside from the ultimate stability or instability of metastable and irregular initial data, several questions remain. For one, black holes formed in massive scalar collapse in asymptotically flat spacetime exhibit a mass gap for initial profiles wider than the Compton wavelength $1/\mu$ [?]. Whether this mass gap exists in AdS is not clear, and it may disappear through repeated gravitational focusing as the field oscillates many times across AdS; investigating this type of critical behavior will likely require techniques similar to those of [?]. Returning to our stability phase diagram, the physical mechanism responsible for chaos that seems to occur for some irregular initial data is not yet clear. Is it some generalization of the same mechanism as found in the two-shell system? Also, would an alternate definition of approximate thermalization in the dual CFT, such as development of a power-law spectrum, lead to a different picture of the stability phase diagram? Finally, the big question is whether there is some test that could be performed on initial data alone that would predict in advance its behavior? So far, no test is entirely successful, so new ideas are necessary.

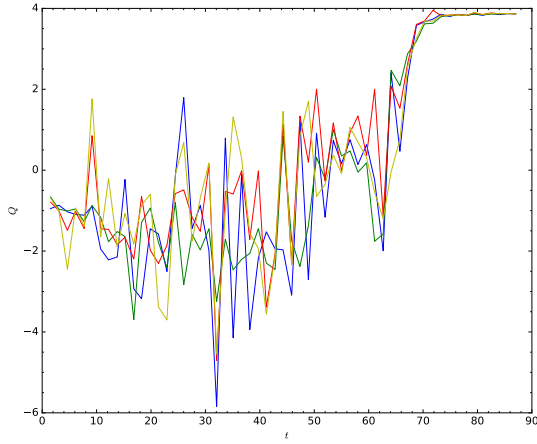
Acknowledgments We would like to thank Brayden Yarish for help submitting jobs for the $\mu = 10$ evolutions. The work of ND is supported in part by a Natural Sciences and Engineering Research Council of Canada PGS-D grant to ND, NSF Grant PHY-1606654 at Cornell University, and by a grant from the Sherman Fairchild Foundation. The work of BC and AF is supported by the Natural Sciences and Engineering Research Council of Canada Discovery Grant program. This research was enabled in part by support provided by WestGrid (www.westgrid.ca) and Compute Canada Calcul Canada (www.computecanada.ca).

Appendix

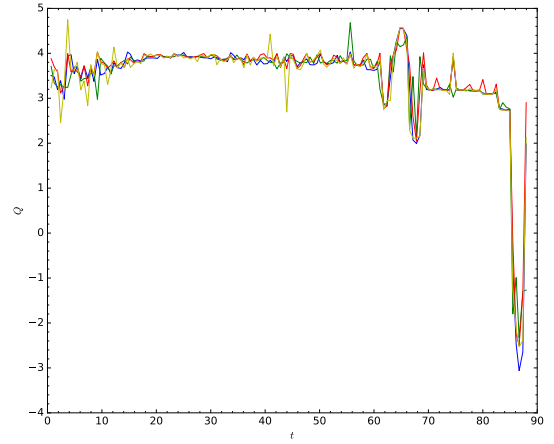
1.A Convergence Testing

Due to the large number of evolutions we have carried out, it is not computationally feasible to test all of them for convergence. Therefore, we have checked several interesting cases of irregular initial data, which are the most curious. These are carried out by evolving the initial data with a base resolution $n = 14$ and again at $n = 15, 16$ with commensurate time steps, as described in [?]. In the cases indicated, we evaluated the order of convergence at lower resolutions. We remind the reader that the order of convergence Q is the base-2 logarithm of the ratio of L^2 errors (root-mean-square over all corresponding grid points) between successive pairs of resolutions. We also note that the initial data is defined analytically, so Q can appear poor at $t = 0$ since the errors are controlled by round off; in some cases, Q is therefore undefined and not plotted.

First, we carried out convergence tests for mass $\mu = 0.5$, width $\sigma = 1$, and amplitude $\epsilon = 1.12$, which is monotonic irregular initial data presented in figure 1.5a. This amplitude collapses with $t_H \sim 88$. Figure 1.A.1a shows the (L^2 norm) order of convergence for the field variable ϕ , the mass function M , and the metric functions A, δ . While the order of convergence is initially poor and even negative, all these variables show approximately fourth order convergence for times $t \gtrsim 70$. The reason for the initially poor convergence is that the error between successive resolutions is already given by (machine limited) round off. As a demonstration, we tested the order of convergence with base resolution $n = 11$, as shown in figure 1.A.1b. The variables show order of convergence $Q \gtrsim 3$ already at this resolution for most of the evolution, losing convergence only for $t > 80$, where we



(a) $n = 14$ base resolution



(b) $n = 11$ base resolution

Figure 1.A.1: Convergence results for $\mu = 0.5$, $\sigma = 1$, $\epsilon = 1.12$ showing order of convergence Q vs time for ϕ, M, A, δ (blue, green, red, yellow respectively). Left: Resolutions $n = 14, 15, 16$ used. Right: Resolutions $n = 11, 12, 13$ used.

see approximately 4th-order convergence in the $n = 14$ resolution computations.

Two of the authors have discussed the convergence properties of evolution for the nonmonotonic irregular initial data with $\mu = 20, \sigma = 0.1, \epsilon = 11.74$, which is in an amplitude region of increased t_H surrounded by smaller values, in detail in [?]. In short, the variables ϕ, M, A, δ all exhibit fourth order convergence, as does $\Pi^2(t, 0)$, and the conserved mass actually has 6th order convergence.

Initial data for $\mu = 15, \sigma = 0.2$ is also nonmonotonic, as shown in figure 1.A.2a. While we have not analyzed all aspects of the convergence, we see from the remainder of figure 1.A.2 that ϕ, M, A, δ exhibit convergent behavior at better than second order for $\epsilon = 7.42$ (figure 1.A.2b, second-largest value of t_H in figure 1.A.2a) and $\epsilon = 7.40$ (figure 1.A.2c, adjacent amplitude in figure 1.A.2a). It is important to note that the larger amplitude also has the larger horizon formation time, contrary to the usual monotonic behavior. In other words, we have validated the nonmonotonicity of this initial data through convergence testing.

It is most crucial to validate the convergence of chaotic evolutions. In table 1.2, we noted that

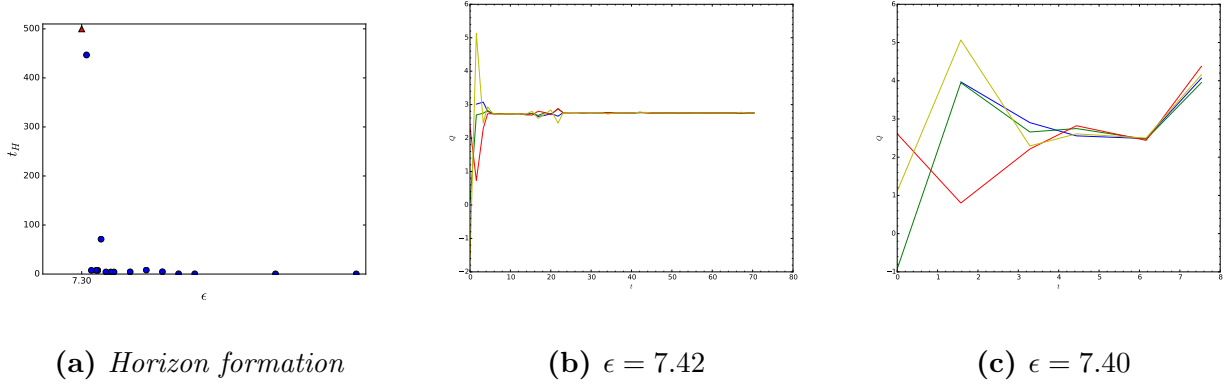


Figure 1.A.2: Convergence results for $\mu = 15$, $\sigma = 0.2$. Left: t_H vs ϵ . Middle & Right: order of convergence vs time for ϕ, M, A, δ (blue, green, red, yellow respectively) for indicated amplitudes.

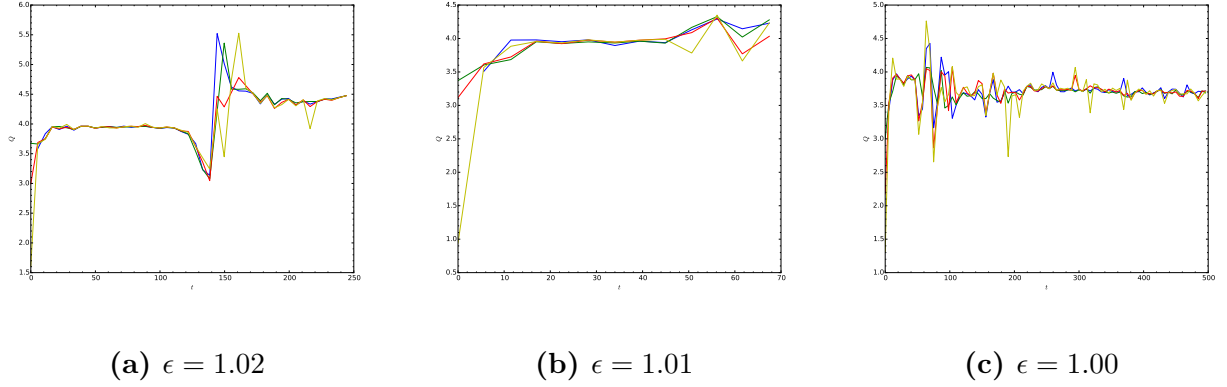


Figure 1.A.3: Convergence results for $\mu = 0$, $\sigma = 1.1$ for listed amplitudes showing order of convergence Q vs time for ϕ, M, A, δ (blue, green, red, yellow respectively); resolutions $n = 12, 13, 14$.

the Ricci scalar at the origin has nonzero Lyapunov exponent at almost the 2 sigma level for amplitudes $\epsilon = 1.02, 1.01, 1.00$ for $\mu = 0, \sigma = 1.1$. We show the results of convergence tests for these amplitudes in figure 1.A.3; because these are longer evolutions, we consider the convergence at the lower resolutions $n = 12, 13, 14$. After a transient start-up period, these are all convergent with $Q > 2.5$ for all variables considered at all times; for most of the time, the order of convergence is $Q > 3.5$. It is worth noting that one of the amplitudes does not form a horizon through $t = 500$. These convergence tests validate both the nonmonotonic nature of the evolution ($t_H \approx 248, 71$ and > 500 for $\epsilon = 1.02, 1.01, 1.00$ respectively) and also the calculation of the Lyapunov coefficient.

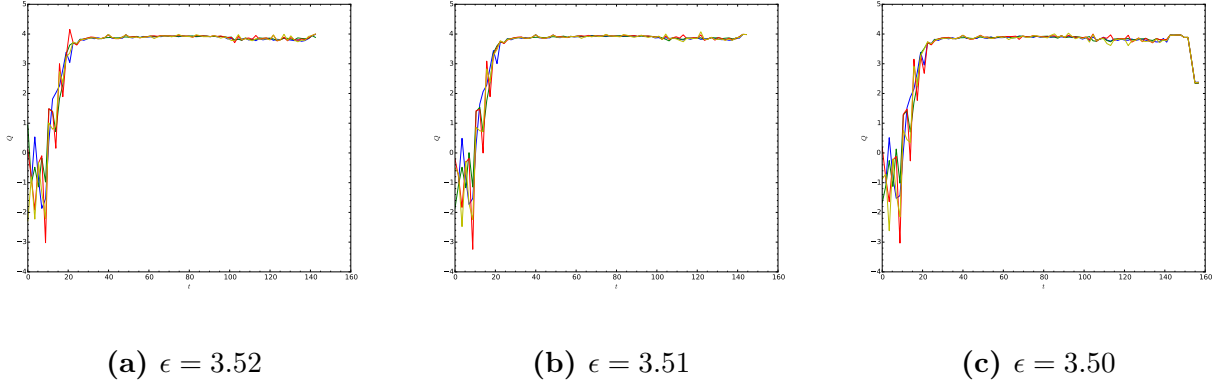


Figure 1.A.4: Convergence results for $\mu = 5$, $\sigma = 0.34$ for listed amplitudes showing order of convergence Q vs time for ϕ, M, A, δ (blue, green, red, yellow respectively); resolutions $n = 14, 15, 16$.

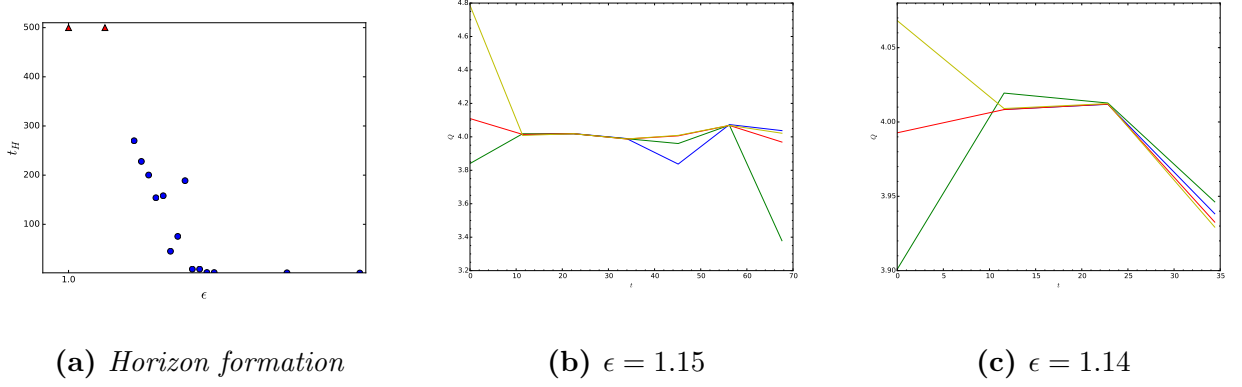
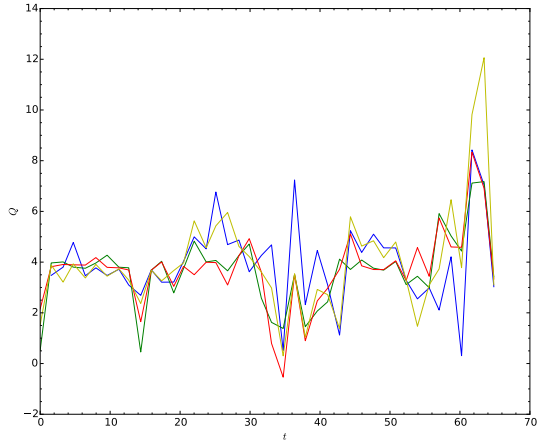


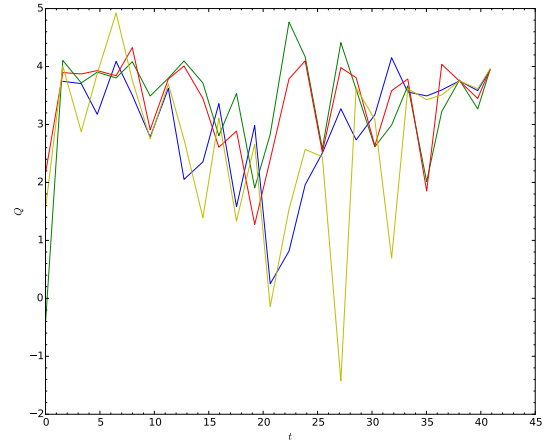
Figure 1.A.5: Convergence results for $\mu = 1$, $\sigma = 1$. Left: t_H vs ϵ . Middle & Right: order of convergence Q vs time for ϕ, M, A, δ (blue, green, red, yellow respectively); resolutions $n = 11, 12, 13$.

Also in table 1.2, we found a nonzero Lyapunov exponent for $\mu = 5, \sigma = 0.34$ at amplitudes $\epsilon = 3.52, 3.51, 3.50$. The results of convergence tests for these amplitudes appear in figure 1.A.4. For $t \gtrsim 20$, these evolutions exhibit convergent behavior with $Q > 3.5$ (and always $Q > 2$). At early times, the apparent poor convergence is again due to the errors being dominated by round-off; we have carried out additional convergence tests (not shown) and verified that these evolutions are already convergent with order of convergence close to $Q = 4$ at base resolutions $n = 12$ for $t \lesssim 20$. Again, convergence tests validate chaotic behavior for these initial data.

Initial data with $\mu = 1, \sigma = 1$ is chaotic over a narrow range of amplitudes. We have carried out



(a) $\epsilon = 6.95$



(b) $\epsilon = 6.92$

Figure 1.A.6: *Order of convergence vs time for ϕ, M, A, δ (blue, green, red, yellow) for $\mu = 20, \sigma = 0.19$ and indicated amplitudes.*

convergence testing for amplitudes $\epsilon = 1.15, 1.14$, which are the two amplitudes with $t_H < 100$ between amplitudes with $t_H \gtrsim 150$ in figure 1.A.5a. The order of convergence was poor for these amplitudes in our initial tests with base resolution $n = 14$ because the error between resolutions was dominated by round-off, similar to the convergence tests we discussed above for $\mu = 0.5, \sigma = 1$. In subsequent tests with lower resolutions $n = 11, 12, 13$, we find an order of convergence $Q \sim 4$ for most of the evolutions (and always $Q > 3$). It is important to note again that our evolutions exhibit convergence while showing horizon formation at a later time for a larger amplitude in this case, again validating the nonmonotonic behavior.

Finally, we ran convergence tests for the chaotic initial data with $\mu = 20, \sigma = 0.19$ for $\epsilon = 6.95, 6.92$, with $t_H \approx 65.5, 40.8$ respectively. As shown in figure 1.A.6, the simulations are close to fourth order convergence for most of the evolution, but there are periods where the order of convergence for evolution and constraint variables becomes negative. This of course leads to the concern that the evolutions should have collapsed during those periods and extend into an “afterlife” evolution. We have therefore evolved these amplitudes through these regions (approximately $t = 30 - 40$ for $\epsilon = 6.95$ and $t = 18 - 30$ for $\epsilon = 6.92$) at high resolution ($n = 18$). If the evolutions are truly in

an afterlife, this higher resolution calculation may include horizon formation. We do not observe this. Another tell-tale of would-be horizon formation is a decrease in the timestep size by an order of magnitude or more followed by an increase. We monitor the timestep size every 500 time steps through this evolution but do not observe a decrease in timestep size by more than a factor of 2. As a result, we believe the values of t_H found are reliable, though the reader may wish to consider them with some caution. In other words, while convergence testing is the gold standard to validate our numerical evolutions, there are other indicators of reliability, which these evolutions satisfy. It is also worth noting that the rapid energy transfer characteristic of figure 1.11d for $\epsilon = 6.95$ begins immediately and is therefore seen in a convergent region of the evolutions, particularly for $t \lesssim 14$. Nonetheless, we emphasize that we have found convergent evolutions for irregular initial data at scalar masses from $\mu = 0$ to 20. It is important to note that we have validated nonmonotonic behavior in plots of t_H vs ϵ . Convergence testing also specifically validates the evolutions used to find a nonzero Lyapunov coefficient (at nearly the 2σ level) for massless scalar collapse.

Bibliography

- [1] P. Bizon and A. Rostworowski, *On weakly turbulent instability of anti-de Sitter space*, [*Phys. Rev. Lett.* **107** \(2011\) 031102](#), [[1104.3702](#)].
- [2] D. Garfinkle and L. A. Pando Zayas, *Rapid Thermalization in Field Theory from Gravitational Collapse*, [*Phys. Rev.* **D84** \(2011\) 066006](#), [[1106.2339](#)].
- [3] J. Jalmuzna, A. Rostworowski and P. Bizon, *A Comment on AdS collapse of a scalar field in higher dimensions*, [*Phys. Rev.* **D84** \(2011\) 085021](#), [[1108.4539](#)].
- [4] D. Garfinkle, L. A. Pando Zayas and D. Reichmann, *On Field Theory Thermalization from Gravitational Collapse*, [*JHEP* **02** \(2012\) 119](#), [[1110.5823](#)].
- [5] O. J. C. Dias, G. T. Horowitz and J. E. Santos, *Gravitational Turbulent Instability of Anti-de Sitter Space*, [*Class. Quant. Grav.* **29** \(2012\) 194002](#), [[1109.1825](#)].
- [6] V. Balasubramanian, A. Buchel, S. R. Green, L. Lehner and S. L. Liebling, *Holographic Thermalization, Stability of Anti-de Sitter Space, and the Fermi-Pasta-Ulam Paradox*, [*Phys. Rev. Lett.* **113** \(2014\) 071601](#), [[1403.6471](#)].
- [7] B. Craps, O. Evnin and J. Vanhoof, *Renormalization group, secular term resummation and AdS (in)stability*, [*JHEP* **10** \(2014\) 048](#), [[1407.6273](#)].
- [8] P. Basu, C. Krishnan and A. Saurabh, *A stochasticity threshold in holography and the instability of AdS*, [*Int. J. Mod. Phys.* **A30** \(2015\) 1550128](#), [[1408.0624](#)].
- [9] F. V. Dimitrakopoulos, B. Freivogel, M. Lippert and I.-S. Yang, *Position space analysis of the AdS (in)stability problem*, [*JHEP* **08** \(2015\) 077](#), [[1410.1880](#)].

- [10] A. Buchel, S. R. Green, L. Lehner and S. L. Liebling, *Conserved quantities and dual turbulent cascades in anti-de Sitter spacetime*, *Phys. Rev.* **D91** (2015) 064026, [[1412.4761](#)].
- [11] B. Craps, O. Evnin and J. Vanhoof, *Renormalization, averaging, conservation laws and AdS (in)stability*, *JHEP* **01** (2015) 108, [[1412.3249](#)].
- [12] O. Evnin and C. Krishnan, *A Hidden Symmetry of AdS Resonances*, *Phys. Rev.* **D91** (2015) 126010, [[1502.03749](#)].
- [13] F. Dimitrakopoulos and I.-S. Yang, *Conditionally extended validity of perturbation theory: Persistence of AdS stability islands*, *Phys. Rev.* **D92** (2015) 083013, [[1507.02684](#)].
- [14] S. R. Green, A. Maillard, L. Lehner and S. L. Liebling, *Islands of stability and recurrence times in AdS*, *Phys. Rev.* **D92** (2015) 084001, [[1507.08261](#)].
- [15] B. Craps, O. Evnin and J. Vanhoof, *Ultraviolet asymptotics and singular dynamics of AdS perturbations*, *JHEP* **10** (2015) 079, [[1508.04943](#)].
- [16] B. Craps, O. Evnin, P. Jai-akson and J. Vanhoof, *Ultraviolet asymptotics for quasiperiodic AdS₄ perturbations*, *JHEP* **10** (2015) 080, [[1508.05474](#)].
- [17] B. Craps and O. Evnin, *AdS (in)stability: an analytic approach*, *Fortsch. Phys.* **64** (2016) 336–344, [[1510.07836](#)].
- [18] A. Buchel, S. L. Liebling and L. Lehner, *Boson stars in AdS spacetime*, *Phys. Rev.* **D87** (2013) 123006, [[1304.4166](#)].
- [19] M. Maliborski and A. Rostworowski, *A comment on "Boson stars in AdS"*, [1307.2875](#).
- [20] M. Maliborski and A. Rostworowski, *Lecture Notes on Turbulent Instability of Anti-de Sitter Spacetime*, *Int. J. Mod. Phys.* **A28** (2013) 1340020, [[1308.1235](#)].
- [21] P. Bizoń and A. Rostworowski, *Comment on "Holographic Thermalization, Stability of Anti-de Sitter Space, and the Fermi-Pasta-Ulam Paradox"*, *Phys. Rev. Lett.* **115** (2015) 049101, [[1410.2631](#)].

- [22] V. Balasubramanian, A. Buchel, S. R. Green, L. Lehner and S. L. Liebling, *Reply to Comment on “Holographic Thermalization, Stability of Anti-de Sitter Space, and the Fermi-Pasta-Ulam Paradox”*, *Phys. Rev. Lett.* **115** (2015) 049102, [[1506.07907](#)].
- [23] P. Bizoń, M. Maliborski and A. Rostworowski, *Resonant Dynamics and the Instability of Anti-de Sitter Spacetime*, *Phys. Rev. Lett.* **115** (2015) 081103, [[1506.03519](#)].
- [24] N. Deppe, *On the stability of anti-de Sitter spacetime*, [1606.02712](#).
- [25] F. V. Dimitrakopoulos, B. Freivogel, J. F. Pedraza and I.-S. Yang, *Gauge dependence of the AdS instability problem*, *Phys. Rev.* **D94** (2016) 124008, [[1607.08094](#)].
- [26] F. V. Dimitrakopoulos, B. Freivogel and J. F. Pedraza, *Fast and Slow Coherent Cascades in Anti-de Sitter Spacetime*, *Class. Quant. Grav.* **35** (2018) 125008, [[1612.04758](#)].
- [27] S. L. Liebling and G. Khanna, *Scalar collapse in AdS with an OpenCL open source code*, *Class. Quant. Grav.* **34** (2017) 205012, [[1706.07413](#)].
- [28] P. Bizoń and J. Jałmużna, *Globally regular instability of AdS_3* , *Phys. Rev. Lett.* **111** (2013) 041102, [[1306.0317](#)].
- [29] E. da Silva, E. Lopez, J. Mas and A. Serantes, *Collapse and Revival in Holographic Quenches*, *JHEP* **04** (2015) 038, [[1412.6002](#)].
- [30] N. Deppe, A. Kolly, A. R. Frey and G. Kunstatter, *Black Hole Formation in AdS Einstein-Gauss-Bonnet Gravity*, *JHEP* **10** (2016) 087, [[1608.05402](#)].
- [31] N. Deppe, A. Kolly, A. Frey and G. Kunstatter, *Stability of AdS in Einstein Gauss Bonnet Gravity*, *Phys. Rev. Lett.* **114** (2015) 071102, [[1410.1869](#)].
- [32] R. Arias, J. Mas and A. Serantes, *Stability of charged global AdS_4 spacetimes*, *JHEP* **09** (2016) 024, [[1606.00830](#)].
- [33] H. Okawa, J. C. Lopes and V. Cardoso, *Collapse of massive fields in anti-de Sitter spacetime*, [1504.05203](#).

- [34] N. Deppe and A. R. Frey, *Classes of Stable Initial Data for Massless and Massive Scalars in Anti-de Sitter Spacetime*, *JHEP* **12** (2015) 004, [[1508.02709](#)].
- [35] V. E. Hubeny and M. Rangamani, *Unstable horizons*, *JHEP* **05** (2002) 027, [[hep-th/0202189](#)].
- [36] A. Buchel and L. Lehner, *Small black holes in $AdS_5 \times S^5$* , *Class. Quant. Grav.* **32** (2015) 145003, [[1502.01574](#)].
- [37] A. Buchel, *Universality of small black hole instability in AdS/CFT*, *Int. J. Mod. Phys. D* **26** (2017) 1750140, [[1509.07780](#)].
- [38] A. Buchel and M. Buchel, *On stability of nonthermal states in strongly coupled gauge theories*, [1509.00774](#).
- [39] R. Brito, V. Cardoso and J. V. Rocha, *Interacting shells in AdS spacetime and chaos*, *Phys. Rev. D* **94** (2016) 024003, [[1602.03535](#)].
- [40] O. Aharony, S. S. Gubser, J. M. Maldacena, H. Ooguri and Y. Oz, *Large N field theories, string theory and gravity*, *Phys. Rept.* **323** (2000) 183–386, [[hep-th/9905111](#)].
- [41] H. Nastase, *Introduction to the ADS/CFT Correspondence*. Cambridge University Press, 2015.
- [42] P. R. Brady, C. M. Chambers and S. M. C. V. Goncalves, *Phases of massive scalar field collapse*, *Phys. Rev. D* **56** (1997) 6057–6061, [[gr-qc/9709014](#)].
- [43] D. Santos-Oliván and C. F. Sopuerta, *Moving closer to the collapse of a massless scalar field in spherically symmetric anti-de Sitter spacetimes*, *Phys. Rev. D* **93** (2016) 104002, [[1603.03613](#)].