

UNIVERSITY OF OSLO

AST3310

TERM PROJECT 3

Title

Written by:
Bendik SAMSETH

April 24, 2016

Contents

1	Introduction	2
2	Adding Convection	2
2.1	Deriving the new equation	2
2.2	Instability criterion	3

Abstract

In this report. . . . All source material related to this report, which is cited at the relevant points, can be found at the projects GitHub repository [1].

1 Introduction

Hey hey, making a star, huh?

2 Adding Convection

Following from the last report, we know that the governing equations for a star is:

$$\frac{\partial r}{\partial m} = \frac{1}{4\pi r^2 \rho} \quad (1)$$

$$\frac{\partial P}{\partial m} = -\frac{Gm}{4\pi r^4} \quad (2)$$

$$\frac{\partial L}{\partial m} = \epsilon \quad (3)$$

$$\frac{\partial T}{\partial m} = -\frac{3\kappa L}{256\pi^2 \sigma r^4 T^3} \quad (4)$$

Adding convective energy transport equates to changing the last equation. This is because the temperature gradient in Eq. (4) is based on all the energy being moved out of the star through radiation. When convection kicks in, the temperature gradient will become smaller, because not all the energy needs to be transferred by radiation anymore.

2.1 Deriving the new equation

So, we want to find the new temperature gradient, including both convection and radiation. We still know that the total energy flux must be related to the luminosity:

$$F_R + F_C = \frac{L}{4\pi r^2} \quad (5)$$

We still have the radiative flux, given as

$$F_R = \frac{4acGmT^4}{3\kappa Pr^2} \nabla \quad (6)$$

We get Eq. (4) by setting $F_C = 0$ and inserting into Eq. (5).

This time, we are not setting $F_C = 0$, obviously, so what should we set it to? The full derivation is given in [2], and is too involved to be repeated in this report. A quick summary will do.

We start from first principles, where we can say that $F_C = \rho c_P v \Delta T$, which says that the

radiative flux is the energy $\rho c_P \Delta T$ a parcel of gas with density ρ can move, with velocity v . The hard part is finding v and the temperature difference. To this end, we will define three quantities, a shorthand symbol δ , the pressure scale height H_P ¹, and temperature gradient ∇ :

$$\delta = \frac{T}{V} \left(\frac{\partial V}{\partial T} \right)_P \quad (7)$$

$$H_P = -P \frac{\partial r}{\partial P} = \frac{Pr^2}{\rho Gm} \quad (8)$$

$$\nabla = \frac{\partial \ln T}{\partial \ln P} = -\frac{H_P}{T} \frac{\partial T}{\partial r}. \quad (9)$$

We want an equation for ∇ , as this would give us the temperature gradient we seek.

It can be shown (derivation once again in [2]) that

$$v = \sqrt{\frac{g \delta l_m^2}{4H_P}} (\nabla - \nabla^*)^{1/2} \quad (10)$$

$$F_C = \rho c_P T \sqrt{g \delta H_P^{-\frac{3}{2}}} \left(\frac{l_m}{2} \right)^2 (\nabla - \nabla^*)^{3/2} \quad (11)$$

where l_m is a constant with unit length. ∇^* is the temperature gradient for the little parcel of gas moving through the surroundings in the star. We only care about the actual ∇ , so we want to eliminate ∇^* . But, before we can do that, we introduce two more ∇ 's, ∇_{ad} and ∇_{rad} . These represent the temperature gradient if we had an adiabatic expansion of the parcel, and the temperature gradient needed to transport all the energy through radiation, respectively.

We can say that $F_C + F_R = F_{rad}$, meaning that the total flux from convection *and* radiation is equal to the flux we would have if all energy to be moved by radiation (i.e. ∇_{rad}). Let's insert our expressions.

¹We get the expression for H_P by using the assumption of hydrostatic equilibrium.

$$\begin{aligned}
F_R + F_C &= \frac{4acGmT^4}{3\kappa Pr^2} \nabla_{rad} \\
F_C &= \frac{4acGmT^4}{3\kappa Pr^2} (\nabla_{rad} - \nabla) \\
(\nabla - \nabla^*)^{3/2} &= \frac{4acGmT^4}{3\kappa Pr^2} \frac{(\nabla_{rad} - \nabla)}{\rho c_P T \sqrt{g\delta} H_P^{-\frac{3}{2}}} \\
(\nabla - \nabla^*)^{3/2} &= \frac{64\sigma T^3}{3\kappa \rho^2 c_P l_m^2} \sqrt{\frac{H_P}{g\delta}} (\nabla_{rad} - \nabla) \\
&\equiv \frac{U}{l_m^2} (\nabla_{rad} - \nabla). \tag{12}
\end{aligned}$$

where we have defined for convenience the constant U (which represents an energy in this case):

$$U = \frac{64\sigma T^3}{3\kappa \rho^2 c_P} \sqrt{\frac{H_P}{g\delta}}. \tag{13}$$

To continue, we need to use one more magic formula[2, ??],

$$\begin{aligned}
(\nabla^* - \nabla_{ad}) &= \frac{32\sigma T^3}{3\kappa \rho^2 c_P v} \frac{S}{dQ} (\nabla - \nabla^*) \\
&\equiv \frac{32\sigma T^3}{3\kappa \rho^2 c_P v} l_m K (\nabla - \nabla^*), \tag{14}
\end{aligned}$$

where $K = S/l_m dQ$ is a geometric constant defined for convenience. If we assume that the parcel is a sphere with diameter $d = l_m$, the parcel has a cross section Q and surface area S , and K reduces to

$$K = \frac{S}{l_m dQ} = \frac{4\pi r_p^2}{l_m^2 \pi r_p^2} = \frac{4}{l_m^2}. \tag{15}$$

Now we insert Eq. (14), combined with Eq. (10) into the following equation:

$$\begin{aligned}
(\nabla^* - \nabla_{ad}) &= (\nabla - \nabla_{ad}) - (\nabla - \nabla^*) \\
\frac{32\sigma T^3}{3\kappa \rho^2 c_P v} l_m K (\nabla - \nabla^*) &= (\nabla - \nabla_{ad}) - (\nabla - \nabla^*) \\
(\nabla - \nabla^*) \left[UK(\nabla - \nabla^*)^{-1/2} + 1 \right] &= \nabla - \nabla_{ad} \\
\Rightarrow \xi^2 + UK\xi - (\nabla - \nabla_{ad}) &= 0 \\
\Rightarrow \nabla &= \xi^2 + UK\xi + \nabla_{ad}
\end{aligned}$$

where $\xi = (\nabla - \nabla^*)^{1/2}$.

What we have is now an expression for ∇ , given by ξ and ∇_{ad} . The latter we know, so we need to find ξ . We do this by inserting ξ into Eq. (12):

$$\begin{aligned}
(\nabla - \nabla^*)^{3/2} &= \xi^3 = \frac{U}{l_m^2} (\nabla_{rad} - \nabla) \\
\xi^3 &= \frac{U}{l_m^2} (\nabla_{rad} - (\xi^2 + UK\xi + \nabla_{ad})) \\
\Rightarrow \xi^3 + \frac{U}{l_m^2} (\xi^2 + UK\xi - (\nabla_{rad} - \nabla_{ad})) &= 0. \tag{16}
\end{aligned}$$

Eq. (16) is a third order equation in ξ , but all the other parameters is known, so we can in principle solve this equation, and thereby getting an expression for the true temperature gradient ∇ .

So, the steps we need to take are:

1. Solve Eq. (16) to get ξ .
2. Insert into $\nabla = \xi^2 + UK\xi + \nabla_{ad}$.
3. Solve for the gradient with respect to mass:

$$\nabla = \frac{P}{T} \frac{\partial T}{\partial P} = \frac{P}{T} \frac{\partial m}{\partial P} \frac{\partial T}{\partial m} \tag{17}$$

$$= \frac{P}{T} \left(-\frac{4\pi r^4}{Gm} \right) \frac{\partial T}{\partial m} \tag{18}$$

$$\Rightarrow \frac{\partial T}{\partial m} = -\frac{GmT}{4\pi r^4 P} \nabla. \tag{19}$$

2.2 Instability criterion

For our code, we would like to know if convection is taking place (and we need to use the harder formula derived in the last section), or if we can just use the old Eq. (4).

Convection will start if the temperature gradient $\partial T/\partial r$ is not large enough to transport all the energy by radiation. We say this by the inequality $\frac{\partial T}{\partial r} < \left(\frac{\partial T}{\partial r} \right)_{ad}$, or equivalently

$$\nabla_{rad} > \nabla_{ad}. \tag{20}$$

So, the code will check if this is the case, and if so, apply the convection formula.

It should be noted that if this criteria is in fact not the case, and $\nabla_{rad} = \nabla_{ad}$, then ξ will have the trivial solution of $\xi = 0$, and consequently we end up with just the same equation as in Eq. (4). It is, however, still a good idea to skip the calculations if they are not necessary, so we keep the two cases separated.

References

- [1] B. Samseth. GitHub repository. URL: <https://github.com/bsamseth/ast3310>.
- [2] B. V. Gudiksen. *AST3310: Astrophysical plasma and stellar interiors*. 2016. URL: <http://www.uio.no/studier/emner/matnat/astro/AST3310/v16/noter/notes-main.pdf> (visited on 04/21/2016).