

# A MULTI-TELESCOPE CAMPAIGN ON FRB 121102: IMPLICATIONS FOR THE FRB POPULATION

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## ABSTRACT

We present results of the coordinated observing campaign that made the first subarcsecond localization of a Fast Radio Burst, FRB 121102. Of the nine bursts detected by the Very Large Array at 3 GHz, three had simultaneous observing coverage at other observatories. Only one VLA burst was detected by another observatory, the Arecibo Observatory at 1.4 GHz. We use multi-observatory constraints and modeling of bursts seen only at 3 GHz to confirm earlier results showing that burst spectra are not well characterized by a power law. We find that burst spectra are characterized by a  $\sim 500$  MHz envelope and apparent radio energy as high as  $10^{40}$  erg. We attribute significant changes in the apparent dispersion between bursts to frequency-dependent profiles that add a systematic random to the estimate of DM by up to 1%. We use FRB 121102 as a prototype of the FRB class to estimate a volumetric birth rate of FRB sources  $R_{\text{FRB}} \approx 5 \times 10^{-5}/N_r \text{ Mpc}^{-3} \text{ yr}^{-1}$ , where  $N_r$  is the number of bursts per source over its lifetime. This rate is broadly consistent with models of FRBs from young pulsars or magnetars born in superluminous supernovae or long gamma-ray bursts, if the typical FRB repeats on the order of  $10^3$  times during its lifetime.

## 1. INTRODUCTION

Fast Radio Bursts (FRBs) were discovered ten years ago with the detection of a millisecond-duration radio transient with an anomalously high dispersion measure (Lorimer et al. 2007). The largest measured DMs imply that FRBs originate outside of our Galaxy, potentially at cosmological distances, and are orders of magnitude more luminous than pulses from Galactic pulsars (Thornton et al. 2013). There are now 21 FRBs publicly known<sup>1</sup> with dispersion measures (DMs) as high as  $1600 \text{ pc cm}^{-3}$  and intrinsic temporal widths of order milliseconds. Both their energetics and distance have inspired a wide variety of models and astrophysical applications (e.g., McQuinn 2014; Kulkarni et al. 2014; Fuller & Ott 2015; Connor et al. 2016b; Cordes & Wasserman 2016; Popov & Pshirkov 2016). However, that potential was limited by the lack of a definitive association of an FRB to an extragalactic host.

This paper is part of a series based on the first localization of an FRB and its unambiguous association to an extragalactic host (Chatterjee et al. 2017; Tendulkar et al. 2017; Marcote et al. 2017). FRB 121102, also known as the “repeating FRB”, was discovered (Spitler et al. 2014) in data acquired in 2012 as part of the PALFA survey of the Arecibo Observatory (Cordes et al. 2006). In mid 2015, new Arecibo observations revealed a series of bursts at a similar DM and sky position that rules out cataclysmic models for this source (Spitler et al. 2016). The typical DM measured by early Arecibo observations was  $559 \text{ pc cm}^{-3}$  (Scholz et al. 2016), although somewhat higher values have been seen in more recent observations ( $560.5 \text{ pc cm}^{-3}$ ; Hessels et al., in prep).

Beginning in August of 2015 (MJD 57623), we made the first of nine detections of FRB 121102 with the Karl G. Jansky Very Large Array (VLA; Chatterjee et al. 2017) and localized it with a precision of  $0.1''$ . Deep optical observing with the Gemini Observatory associated the FRB with a host galaxy at  $z = 0.193$  (Tendulkar et al. 2017) and the European VLBI Network detected four more bursts to localize the source with a precision of  $0.01''$ , four orders of magnitude better than any other FRB (precision of  $\sim 40 \text{ pc}$  in linear distance; Marcote et al. 2017). The lookback and luminosity distances for FRB 121102 are 746 and 972 Mpc, respectively, assuming a concordance cosmology with parameters given by Planck Collaboration et al. (2016). If FRB 121102 is representative of all FRBs, then we should expect them to be useful as probes of the intergalactic medium and

their host galaxies. Thus, the confirmation of a cosmological distance for FRB 121102 has begun to fulfil the promise implied by the first FRB detection.

Many new models for the origin of FRBs have been developed in response to the localization of FRB 121102 (Kashiyama & Murase 2017; Metzger et al. 2017; Zhang 2017; Dokuchaev & Eroshenko 2017; Beloborodov 2017; Thompson 2017). The association of FRB 121102 with a compact, persistent radio source is consistent with bursts coming from a young magnetar that powers a luminous pulsar wind nebula (Kashiyama & Murase 2017). At the same time, FRB 121102 is found in a low-metallicity dwarf galaxy (Tendulkar et al. 2017; Metzger et al. 2017), which is not a prediction of models that scale with stellar mass or star formation (Nicholl et al. 2017). These galaxies are the preferred environment for long GRBs and Type-I superluminous supernovae (LGRB and SLSN-I, respectively), which have been suggested are signatures of the magnetar birth (Modjaz et al. 2008; Lunnan et al. 2014).

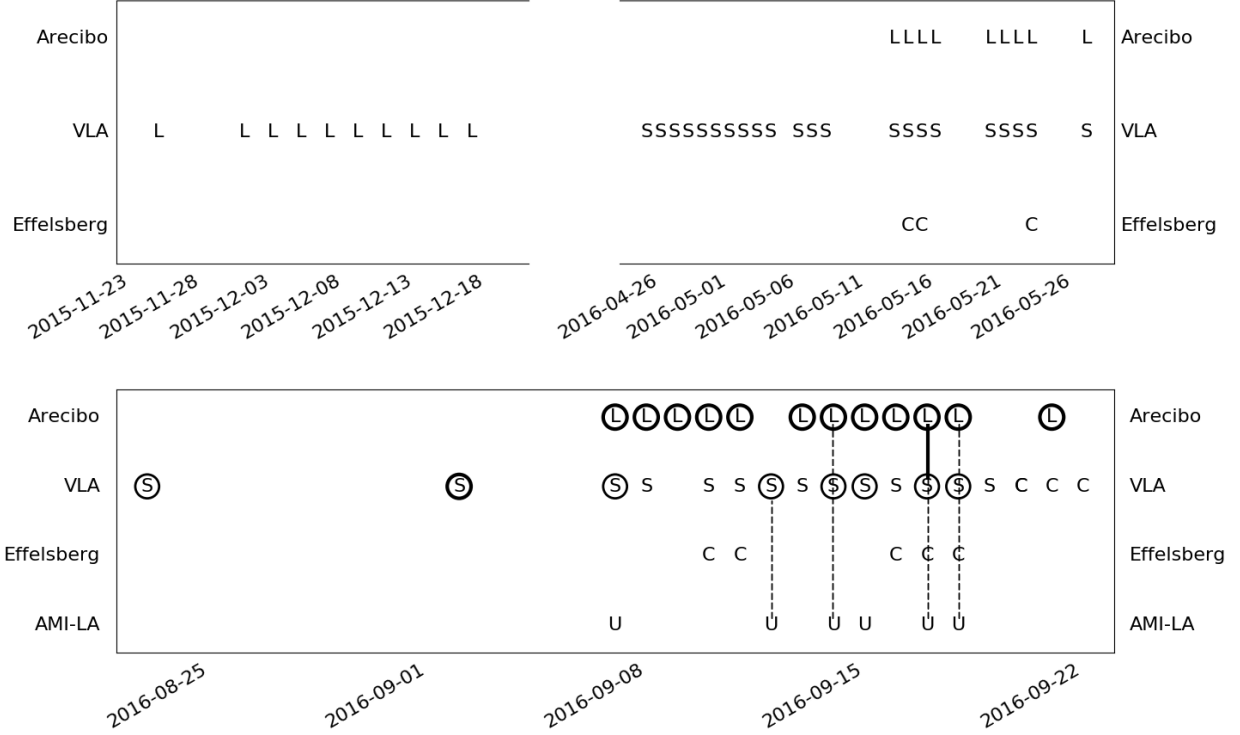
This paper presents an analysis of the spectral properties of VLA bursts implied by simultaneous observing at Effelsberg, GBT, AMI-LA, and Arecibo, including the first simultaneous detection of an FRB at two observatories. If we assume that FRB 121102 is representative of all FRBs, we can use it to constrain the physical processes at play in the overall FRB population. The repetition of FRB 121102 also has strong implications for calculations of their rate of occurrence (Connor et al. 2016a) and comparison to other classes of transient, such as superluminous supernovae (Tendulkar et al. 2017).

In Section 2, we describe the multi-telescope observing campaign and a refined analysis of the nine VLA bursts. Section 3 presents the first spectrum of an FRB simultaneously detected at multiple telescopes and confirm that burst spectra cannot be modeled with a single spectral index. We then model the dynamic spectra to characterize the burst spectra, energies, and dispersion properties. Section 4 discusses the properties of FRB 121102 bursts and its impact on inferences about the birth rate of FRB sources and strategies to find and/or localize new FRBs.

## 2. OBSERVATIONS

The data presented here were obtained from multiple programs and telescopes with a goal of interferometrically localizing FRB 121102 with the VLA. We coordinated observing between the VLA, Arecibo, Effelsberg, and AMI-LA telescopes, as shown in Figure 1 and Table 1. Below, we summarize these observations, with a focus on those conducted simultaneously with VLA burst detections from FRB 121102.

<sup>1</sup> See <http://www.astronomy.swin.edu.au/pulsar/frbcat>.



**Figure 1.** The top and bottom panels summarize the observations and detections of FRB 121102 in the 2015 and 2016 campaigns. Circles highlight observations that detected bursts from FRB 121102 with multiple burst detections indicated with a heavy line. The black dashed lines show the VLA burst detections with simultaneous coverage at other telescopes and the solid line shows the simultaneous burst detection. Circled observations with no line indicate observations on the same day that did not overlap. Symbols show days with observations with letters L, S, C, U referring to radio frequencies of 1.4, 3, 4.5, and 15 GHz, respectively.

Computational (Jupyter) notebooks to reproduce the transient detection, localization, and analysis presented here can be found at <https://github.com/caseyjlaw/FRB121102>. Time cut-out visibility data and calibration products are available at <https://doi.org/10.7910/DVN/TLDKXG>. Original VLA visibility data are available under NRAO program codes 16A-459 and 16A-496 and can be downloaded at <http://archive.nrao.edu>.

**Table 1.** Table of VLA Observations

| Day<br>(yyyy-mm-dd) | Start time<br>(hh:mm:ss) | Duration<br>(min) | Freq.<br>(GHz) | Bursts |
|---------------------|--------------------------|-------------------|----------------|--------|
| 2015-11-25          | 3:25:29                  | 60                | 1.4            | 0      |
| 2015-12-01          | 5:31:31                  | 60                | 1.4            | 0      |
| 2015-12-03          | 2:53:57                  | 60                | 1.4            | 0      |
| 2015-12-05          | 4:45:44                  | 60                | 1.4            | 0      |
| 2015-12-07          | 4:37:50                  | 60                | 1.4            | 0      |

*Table 1 continued*

**Table 1** (*continued*)

|            |          |     |     |   |
|------------|----------|-----|-----|---|
| 2015-12-09 | 9:29:25  | 60  | 1.4 | 0 |
| 2015-12-11 | 9:22:27  | 60  | 1.4 | 0 |
| 2015-12-13 | 9:13:7   | 60  | 1.4 | 0 |
| 2015-12-15 | 9:6:16   | 60  | 1.4 | 0 |
| 2015-12-17 | 8:57:51  | 60  | 1.4 | 0 |
| 2016-04-25 | 15:39:33 | 42  | 3   | 0 |
| 2016-04-26 | 21:49:35 | 15  | 3   | 0 |
| 2016-04-30 | 19:14:46 | 60  | 3   | 0 |
| 2016-05-01 | 18:8:21  | 60  | 3   | 0 |
| 2016-05-03 | 22:53:26 | 60  | 3   | 0 |
| 2016-05-04 | 22:50:30 | 60  | 3   | 0 |
| 2016-05-06 | 18:25:12 | 60  | 3   | 0 |
| 2016-05-07 | 19:34:13 | 60  | 3   | 0 |
| 2016-05-08 | 19:30:20 | 60  | 3   | 0 |
| 2016-05-13 | 17:12:48 | 135 | 3   | 0 |

*Table 1 continued*

**Table 1** (*continued*)

|            |          |     |   |                      |
|------------|----------|-----|---|----------------------|
| 2016-05-14 | 17:14:11 | 135 | 3 | 0                    |
| 2016-05-15 | 17:15:26 | 135 | 3 | 0                    |
| 2016-05-16 | 16:59:52 | 135 | 3 | 0                    |
| 2016-05-20 | 16:45:46 | 135 | 3 | 0                    |
| 2016-05-21 | 16:44:6  | 135 | 3 | 0                    |
| 2016-05-22 | 16:43:53 | 135 | 3 | 0                    |
| 2016-05-23 | 16:44:11 | 135 | 3 | 0                    |
| 2016-05-27 | 16:29:34 | 135 | 3 | 0                    |
| 2016-08-23 | 17:26:28 | 54  | 3 | 1                    |
| 2016-09-02 | 15:52:17 | 54  | 3 | 2                    |
| 2016-09-07 | 10:14:50 | 120 | 3 | 1                    |
| 2016-09-08 | 10:14:40 | 120 | 3 | 0                    |
| 2016-09-10 | 9:59:36  | 120 | 3 | 0                    |
| 2016-09-11 | 9:59:49  | 120 | 3 | 0                    |
| 2016-09-12 | 9:15:19  | 120 | 3 | 1 <sup>a</sup>       |
| 2016-09-13 | 9:23:59  | 120 | 3 | 0                    |
| 2016-09-14 | 9:20:23  | 120 | 3 | 1 <sup>a,b</sup>     |
| 2016-09-15 | 9:16:29  | 120 | 3 | 1                    |
| 2016-09-16 | 9:11:23  | 120 | 3 | 0                    |
| 2016-09-17 | 8:59:20  | 120 | 3 | 1 <sup>a,b,*,c</sup> |
| 2016-09-18 | 8:59:27  | 120 | 3 | 1 <sup>a,b,c</sup>   |
| 2016-09-19 | 8:44:32  | 120 | 3 | 0                    |
| 2016-09-20 | 8:44:33  | 120 | 3 | 0                    |
| 2016-09-20 | 17:19:3  | 120 | 6 | 0                    |
| 2016-09-21 | 9:13:57  | 120 | 6 | 0                    |
| 2016-09-22 | 9:12:24  | 120 | 6 | 0                    |

NOTE—Times are in UT. Frequency is the approximate center of the bandwidth.

<sup>a</sup> AMI-LA coverage at 15 GHz

<sup>b</sup> Arecibo coverage at 1.4 GHz

\* Arecibo detection at 1.4 GHz

<sup>c</sup> Effelsberg coverage at 4.5 GHz

## 2.1. VLA

The FRB 121102 observing campaign started in November 2015 with a 10 hr campaign ( $\sim 1$  hr per session) observed at 1.4 GHz in the compact D configuration. In April through May 2016, we conducted a 40 hr campaign ( $\sim 2$  hr per session) at 3 GHz in the C and CnB configurations in coordination with Arecibo (Scholz et al. 2016). We concluded with a 40 hr, coordinated campaign ( $\sim 2$  hr per session) from August

through September 2016 in the B configuration and during the move to the most extended A configuration. In the late-2016 campaign, the first 34 hours of VLA observations were made at 3 GHz, while the last 6 hours were observed at 6 GHz. These observing session times are inclusive of calibration and overhead, which typically amounts to 15% of the total observing time.

All VLA fast-sampled data were observed with 5 ms sampling, 256 channels, and dual-circular polarization (Law et al. 2015). The channel frequency width was set to maximize sensitivity to the known DM of FRB 121102 with the largest total bandwidth. The total bandwidths at L (1.4 GHz), S (3 GHz), and C (6 GHz) bands were 256 MHz, 1024 MHz, and 2048 MHz, respectively. The 3 GHz data were recorded data in eight spectral windows with 32 channels each and had sensitivity of 5 mJy in 5 ms ( $1\sigma$ ).

Observations in August and September were searched by a prototype version of *realfast* (see <http://realfast.io>; Law et al. 2017). *realfast* is a real-time, fast imaging transient search system. The current prototype runs on CPU-based hardware that is normally dedicated to the VLA correlator; for this experiment, it runs the transient search pipeline software called *rtpipe*<sup>2</sup>. Images were formed for each integration with DMs of 0, 546, 556.9, 560, and 565 pc cm<sup>-3</sup> and a time resampling grid of 5, 10, 20, 40, and 80 ms. This DM grid was chosen to maintain 90% sensitivity to the nominal DM range of 540–570 pc cm<sup>-3</sup>. Gain calibration was made from observations of J0555+3948 by an automated system (telcal), which uses phase-only calibration. A flux scale was calculated for each spectral window from an observation of 3C48 and applied to all burst spectra and has an accuracy of about 10%.

Burst detections and localizations were made within 5–10 hours of data being recorded. The transient search starts when data are recorded and proceeds slower than real-time, so we refer to it as “quasi real-time”. For each trial DM, integration, and time scale, we form an image and calculate the S/N ratio for the peak pixel in the dirty image. We empirically identified S/N thresholds of 6.4 and 7.4 as useful to capture data quality statistics and candidates for inspection, respectively. The higher threshold is relatively unlikely to be triggered by thermal noise in this configuration, so *realfast* generates a more detailed (and computationally intensive) candidate visualization that includes an image and spectrum. All visibilities are recorded so detailed analysis, including

<sup>2</sup> See <https://github.com/caseyjlaw/rtpipe> and Law et al. (2015).

improved calibration and localization, can be conducted offline.

## 2.2. Arecibo

During the late-2016 campaign, Arecibo observed with the L-wide receiver using the PUPPI pulsar backend. The observational frequency range was 1.15 to 1.73 GHz and frequency resolution was 1.5625 MHz, which has a typical sensitivity of 2 mJy in 2 ms ( $1\sigma$ ). We recorded total full Stokes polarization intensity spectra with a time resolution of 10.24  $\mu$ s. Each frequency channel was coherently dedispersed to 557 pc cm<sup>-3</sup>, significantly reducing intra-channel dispersion smearing. The full width at half maximum (FWHM) beam size at band center is 3.3'.

In total, 11 Arecibo observations had some simultaneous coverage with the VLA. All of these observations were conducted after the VLA localization, so they were pointing directly at FRB 121102. Three of those observations had simultaneous coverage of VLA bursts and one of those Arecibo observations detected the VLA burst. Overall, there were many more bursts detected at Arecibo than with the VLA, including some Arecibo bursts with simultaneous VLA upper limits. A more detailed analysis of those bursts will be presented elsewhere (Michilli et al, in prep).

## 2.3. Effelsberg

Effelsberg observations were conducted with the S60mm receiver and recorded total intensity spectra with the PFFTS backend in pulsar search mode. The observations had a time resolution of 65.5  $\mu$ s and 128 frequency channels. The receiver has a system equivalent flux density of 18 Jy and a FWHM beam size of 2.4' at 4.85 GHz.

Five Effelsberg observations were made pointing at the known location of FRB 121102 and had some simultaneous coverage with the VLA. Two of these observations were simultaneous with VLA bursts. The nominal receiver bandwidth is from 4.6 to 5.1 GHz, but a configuration error reduced effective bandwidth to 100 MHz centered at 4.85 GHz for both of these observations during VLA bursts. The sensitivity in these observations was about 28 mJy in 2 ms ( $1\sigma$ ), which is two times worse than the full-bandwidth value. No burst was detected in either observation.

## 2.4. AMI-LA

We observed FRB 121102 with the Arcminute MicroKelvin Imager Large Array (AMI-LA; Zwart et al. 2008) for 3 hours each on four epochs starting at MJDs 57638, 57643, 57645, 57646, 57648, and 57649. Obser-

vations were made with the new digital correlator having 4096 channels across a 5 GHz bandwidth between 13–18 GHz with a 1 s integration time. The phase calibrator, J0518+3306, was observed every 12 minutes for about 1.5 minutes. The AMI-LA data were binned to eight 0.625 GHz channels and processed (RFI excision and calibration) with a fully-automated pipeline, AMI-REDUCE (e.g., Perrott et al. 2013). Daily measurements of 3C48 and 3C286 were used for the absolute flux calibration, which is good to about 10%.

We inspected the calibrated visibilities, and did not find any signal above 30 mJy in the 1 s samples at and in the vicinity of the detected bursts. This corresponds to a sensitivity of 15 Jy for an assumed pulse width of 2 ms. Concatenating and imaging the 12 hours of calibrated data with the CASA tasks *concat* and *clean* also does not yield any significant detection at the FRB location. Although the statistical  $3\sigma$  upper limit is 60  $\mu$ Jy, extended mJy-level sources in the field cause side-lobe confusion (the AMI-LA angular resolution is  $\sim 30''$ ), and the actual upper limit is larger. We introduced artificial point sources at the FRB location using the CASA *sm* tool, and found that these sources can be recovered at  $3\sigma$  as long as their peak flux densities are more than  $\sim 100\mu$ Jy. Hence, we place an upper limit of  $100 \pm 10\mu$ Jy on any quiescent or possible radio flaring on  $\sim$ days timescales from the FRB. This limit is similar to the flux density measured by the VLA at 12 GHz for the persistent source (Chatterjee et al. 2017).

# 3. RESULTS

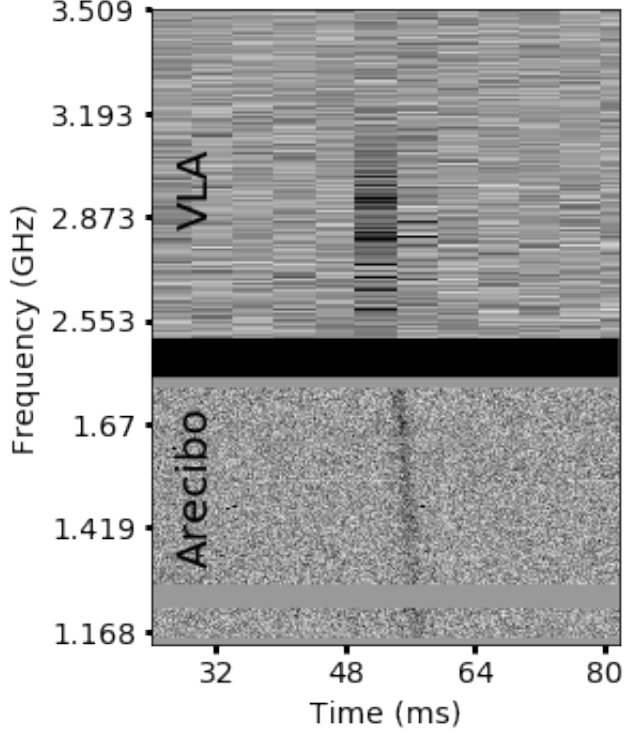
## 3.1. Multi-Observatory Burst Spectrum

Four of the VLA bursts had simultaneous observing coverage with either Arecibo, Effelsberg, or AMI-LA. Two of the VLA bursts had simultaneous observing coverage by all four observatories (Figure 1). Of these bursts, only one VLA burst (on MJD 57648) was simultaneously detected by another observatory (Arecibo).

Figure 2 shows a dynamic spectrum for this burst with both VLA and Arecibo data. The temporal resolution of the VLA data (5 ms) is much coarser than that of Arecibo, so data were resampled to the Arecibo temporal grid after correcting for barycentric and dispersion measure offsets. We assumed a DM of 560.5 pc cm<sup>-3</sup> (Hessels et al., in prep) to calculate a dispersion delay of 189 and 734 ms from infinite frequency to the top of the VLA and Arecibo bands, respectively. We used PRESTO (Ransom 2001) to calculate a relative barycentric time correction of 3.8 ms between the VLA and Arecibo at the time of observation.

The regridded dynamic spectrum has an apparent DM error that is evident both as a frequency-dependent time

drift within the Arecibo observation and as an offset between the Arecibo and VLA bursts. Both of these drifts are consistent with an apparent DM of  $\sim 565 \text{ pc cm}^{-3}$ . More refined modeling of the VLA dynamic spectrum alone (§3.2.1 and Table 2) are consistent with this higher DM value.

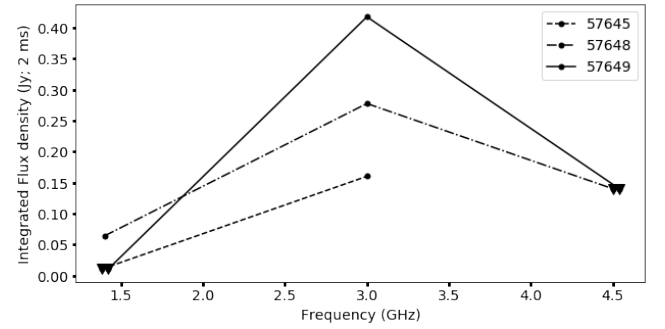


**Figure 2.** Composite dynamic spectrum for the burst from FRB 121102 on MJD 57648 with data from the VLA and Arecibo observatories. The time axis is measured in milliseconds relative to MJD 57648.437788874842 with times corrected to the barycenter and infinite frequency (assuming  $\text{DM} = 560.5 \text{ pc cm}^{-3}$ ). The thick black line separates data from the VLA (2.5 to 3.5 GHz) and Arecibo (1.1 to 1.8 GHz) and flagged data are grey. The amplitude scale is normalized by assuming the Arecibo to VLA gain ratio is 10.

Figure 3 shows coarse spectra built from integrated flux densities measured (or limited) for three bursts with multi-telescope observing coverage. The burst on MJD 57648 was detected with VLA and Arecibo with significances of  $25\sigma$  and  $39\sigma$ , which corresponds to peak flux densities in 5 ms of 111 mJy and 14 mJy at 3 and 1.4 GHz, respectively. Here we assume a typical pulse width of 2 ms to measure VLA and Arecibo integrated flux densities of 278 and 64 mJy at 3 and 1.4 GHz, respectively. Assuming a power law flux density model ( $S_\nu \propto \nu^\alpha$ ), we find a spectral index  $\alpha = 1.9$ . This is inconsistent with the spectral index limit implied by the

Effelsberg nondetection at 4.5 GHz ( $\alpha < -1.7$  for a 2 ms burst and  $5\sigma$  limit).

Overall, the bursts with simultaneous observing coverage are not well described by a power-law model. Two Arecibo nondetections and two Effelsberg nondetections of VLA bursts place mutually inconsistent lower and upper limits on the spectral index. The Arecibo and Effelsberg upper limits at 1.4 and 4.5 GHz, respectively, limit all burst spectral indices  $\alpha_{1.4/3} > +4.8$  and  $\alpha_{3/4.5} < -2.7$ . The strictest limits on  $\alpha_{1.4/3}$  and  $\alpha_{3/4.5}$  are both derived from the burst on MJD 57649, which shows that a power law model is inappropriate even for a single burst.



**Figure 3.** Broadband spectral measurements and limits for three bursts of FRB 121102 with simultaneous multi-telescope observing coverage (Arecibo at 1.4 GHz, VLA at 3 GHz, and Effelsberg at 4.5 GHz). Measurements are shown with dots and upper limits with triangles. Upper limits assume a pulse width of 2 ms and a  $5\sigma$  detection threshold. Overlapping points are offset by tens of MHz in frequency for visualization.

### 3.2. VLA Bursts

#### 3.2.1. Spectral Modeling

This paper refines the analysis of the nine VLA radio burst spectra described in Chatterjee et al. (2017) in a few ways. We use a better calibration scheme and have optimized the detection significance over a fine grid of DM ( $\Delta\text{DM} = 1 \text{ pc cm}^{-3}$ ). After calibration and flagging, the visibility phases were rotated to the best-fit location (RA, Dec = 05h31m58.70s, +33d08m52.5s; Chatterjee et al. 2017) to extract a Stokes I spectrum that maximizes the image S/N for each burst.

Table 2 summarizes the properties of all nine VLA bursts. Parameters such as integrated flux density, S/N, and Stokes V were measured from the burst properties integrated in frequency. Stokes V values estimated from the dual-circular feeds as  $(RR - LL)/(RR + LL) \approx 3\%$ , which is comparable to the systematic error induced by

the primary beam<sup>3</sup>. Given that systematic effects dominate the apparent circular polarization, this observation limits the fractional circular polarization to less than 3%.

Figure 4 shows that the Stokes I spectra are generally characterized by a broad, Gaussian shape with inter-channel modulation as high as 100%. Diffractive scintillation from the Milky Way has a typical bandwidth of 7 MHz along this line of sight (Cordes & Lazio 2002), which is similar to the channel size of 4 MHz. The burst energy is calculated from the flux density in 5 ms and integrating over the optimal Gaussian spectral model.

We used a Markov chain Monte Carlo (MCMC) method to model the dynamic radio spectra (time versus frequency). We used the Goodman and Weare affine invariant sampler (Goodman & Weare 2010) as implemented in *emcee* (Foreman-Mackey et al. 2013). The dynamic spectra were modeled prior to dispersion correction. The spectral structure was modeled as a broad Gaussian envelope (parameters of amplitude, center, width), while the time evolution was modeled by a  $\nu^2$  dispersion law (parameters of starting integration, DM). Finally, the pulses were assumed to have an intrinsically square temporal width (parameter W). This 6-parameter model defined the intrinsic signal that was then distributed over a fixed time-frequency observing grid. The inter-channel amplitude modulation was ignored in the model.

We use a log likelihood function  $\ln \mathcal{L} = (-1/2) \sum_i (S_i - S_{i,\text{model}})^2 / \sigma_s^2$ , where  $S_i$  refers to the measured and modeled flux density in a single pixel of the dynamic spectrum and  $\sigma_s \approx 70 \mu\text{Jy}$  is the measured off-burst noise per 5 ms integration and 4 MHz channel. A flat prior was used over all ranges with valid data with the requirement that the integrated signal must have a detection significance higher than  $8\sigma$ . The 6-dimensional models were sampled with 100 chains taking 700 steps; we ignored the first 200 steps to properly sample the posterior distributions.

The parameters for the best model of the dynamic spectrum are given in Table 2 and the resulting Gaussian model overlaid on Figure 4. The typical burst spectrum has a FWHM of 500 MHz. All but two of the best-fit Gaussians are centered inside the 3 GHz band and some appear contained by the 1 GHz wide band. This is consistent with previous detections of FRB 121102 by Arecibo, which showed quasi-broadband structure

(Scholz et al. 2016) and large variation in the implied spectral index (Spitler et al. 2014).

This analysis differs from traditional analysis in that it simultaneously models the spectral and temporal evolution of the burst. A typical FRB 121102 burst takes  $\sim 180$  ms or 36 integrations to cross the observing band from 2.5 to 3.5 GHz. The effect of a finite pulse width is visible as the burst moves from one integration to the next, even for widths narrower than the integration time of 5 ms. Table 2 shows that the typical burst width is measured as  $2 \pm 0.5$  ms (68% confidence interval); the inter-channel dispersion smearing ranges from 0.4 to 1.1 ms across this band. The brightest burst (on 57633.68) is modeled with a temporal width of  $2.05 \pm 0.02$  ms and its arrival time is measured with a precision of  $\sim 50 \mu\text{s}$ . Note that these errors are only accurate to the degree that the model represents the data. One potential bias in the model is that we do not model scintillation effects.

While most bursts are modeled with well-defined parameter probability distributions, one burst is an outlier. Figure 5 shows the scatter plot matrix (Foreman-Mackey 2016) for the MCMC run on burst 57646. Two clusters of samples are identified for this burst: one narrow, low-DM and one wide, high-DM. Assuming that the model is appropriate, this implies that the burst can be decomposed into two subbursts.

### 3.2.2. Dispersion

The initial VLA detections were made with a matched-filter approach that allowed inter-DM sensitivity losses up to roughly 10% ( $\Delta\text{DM} = 10 \text{ pc cm}^{-3}$  Cordes & McLaughlin 2003). The MCMC modeling is more sophisticated in that it considers a wider range of parameters and models both data and observing effects (frequency-time pixel boundaries) simultaneously.

Figure 6 compares the 68% confidence intervals on the DM for all nine VLA bursts against detection time (top panel) and the modeled burst temporal width (bottom panel). The error weighted mean burst DM is  $567.8 \pm 0.1 \text{ pc cm}^{-3}$ , significantly larger than the long-term average of  $560.5 \text{ pc cm}^{-3}$  seen by Arecibo during this campaign. However, we note that the 95% confidence intervals in DM for multiple bursts are not consistent with this mean nor each other, which suggests that some burst-specific property can bias the measured DM. The variation in DM observed by the VLA at 3 GHz is similar to that reported for Arecibo observations of FRB 121102 in Scholz et al. (2016).

The bottom panel of Figure 6 compares the DM to the modeled temporal width of the bursts. There is a weak correlation between burst width and apparent DM. A

<sup>3</sup> See EVLA memo #195 at [https://library.nrao.edu/public/memos/evla/EVLAM\\_195.pdf](https://library.nrao.edu/public/memos/evla/EVLAM_195.pdf).

**Table 2.** Properties of Bursts from FRB 121102

| Time                   |                     | Observed properties |                              |                | Modeled properties                    |                                     |                              |                 |               |  |
|------------------------|---------------------|---------------------|------------------------------|----------------|---------------------------------------|-------------------------------------|------------------------------|-----------------|---------------|--|
| Calendar day<br>(2016) | Burst time<br>(MJD) | S/N<br>(image)      | $S_{1.5\text{ ms}}$<br>(mJy) | $S_V$<br>(mJy) | DM<br>(pc cm <sup>-3</sup> )          | W<br>(ms)                           | $S_{1,\text{peak}}$<br>(mJy) | Center<br>(GHz) | FWHM<br>(MHz) | $E_{\text{int}}$<br>(10 <sup>38</sup> erg) |
| 23 Aug                 | 57623.74402686      | 38                  | 194                          | +3             | 567±2                                 | 2.0±0.2                             | 690                          | 2.8             | 290           | 11   |
| 2 Sep                  | 57633.67986367      | 179                 | 1500                         | -35            | 568.2±0.2                             | 2.05±0.02                           | 3340                         | 3.2             | 510           | 98   |
| 2 Sep <sup>c</sup>     | 57633.69515938      | 15                  | 69                           | +2             | 562 <sup>+4</sup> <sub>-6</sub>       | 2.5 <sup>+0.9</sup> <sub>-0.6</sub> | >430                         | <2.5            | <290          | 7  |
| 7 Sep                  | 57638.49937435      | 12                  | 55                           | +5             | 567 <sup>+7</sup> <sub>-9</sub>       | 1.3 <sup>+1.4</sup> <sub>-0.8</sub> | 130                          | 3.1             | 420           | 3  |
| 12 Sep <sup>a</sup>    | 57643.45730263      | 100                 | 508                          | -5             | 565.6 <sup>+0.6</sup> <sub>-0.5</sub> | 1.9±0.1                             | 1170                         | 2.8             | 510           | 34   |
| 14 Sep <sup>a</sup>    | 57645.42958602      | 13                  | 64                           | +3             | 563 <sup>+5</sup> <sub>-4</sub>       | 1.1±0.7                             | 170                          | 2.8             | 380           | 4  |
| 15 Sep <sup>c</sup>    | 57646.46600650      | 20                  | 87                           | +1             | 569±5                                 | 2.5 <sup>+0.9</sup> <sub>-1.4</sub> | >420                         | <2.5            | <430          | 10   |
| 17 Sep <sup>a,b</sup>  | 57648.43691490      | 25                  | 111                          | +9             | 564±2                                 | 1.4 <sup>+0.3</sup> <sub>-0.4</sub> | 260                          | 2.8             | 470           | 7  |
| 18 Sep <sup>a</sup>    | 57649.45175697      | 36                  | 167                          | +1             | 567±2                                 | 2.1±0.5                             | 290                          | 3.0             | 690           | 12   |

NOTE— $S_V$  is the measured circular polarization of the burst, which dominated by systematic effects. All error ranges represent 68% confidence intervals.

<sup>a</sup>Simultaneously observed by either Arecibo, Effelsberg, or AMI-LA.

<sup>b</sup>Detected simultaneously with Arecibo.

<sup>c</sup>Best-fit Gaussian is not centered in 3 GHz band, so spectral parameters are limits.

change in width of  $\sim 1$  ms correlates with a change in apparent DM of approximately 5 pc cm<sup>-3</sup>.

### 3.3. Temporal Statistics

The burst rate for FRB 121102 varied dramatically throughout the 2016 observing campaign (Figure 1). In the early-2016 campaign, we observed for 30 hours at 3 GHz and no bursts were detected. In the late-2016 observing campaign, we observed for 27 hours at 3 GHz and detected nine bursts. Overall, the data quality is uniform and high, so the inhomogeneous burst distribution shows that the burst rate was not uniform.

Assuming that the burst detection probability follows a Poisson distribution, the mean VLA, 3 GHz burst rate is  $R = 0.16 \pm 0.05$  hr<sup>-1</sup> above a fluence of 0.2 Jy ms. The nondetection in the first half of 3 GHz observing limits the FRB rate to  $R < 0.1$  hour<sup>-1</sup> (95% confidence limit). The 3 GHz burst rate was much higher during the late-2016 campaign,  $R = 0.3 \pm 0.1$  hour<sup>-1</sup>.

### 3.4. Energy and Brightness Distribution

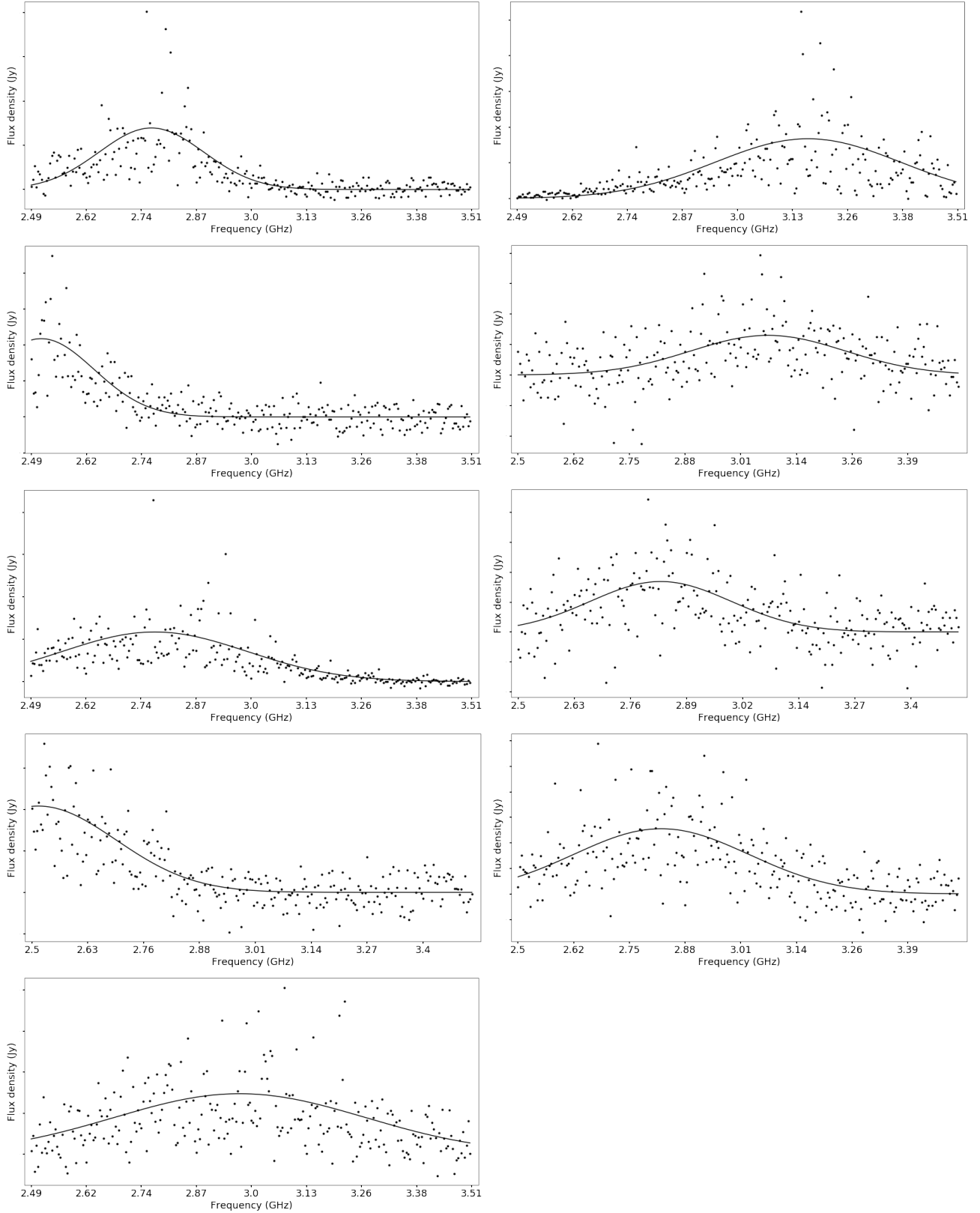
Knowing the source distance, we calculate the burst energy by integrating flux in time and frequency (Table 2). Some VLA burst spectra seem to be contained by the 2.5 to 3.5 GHz band and most of them seem to have Gaussian envelopes that are well modeled by the emission within that band. Assuming that the Gaussian shape defines the full emission window, the mean 3 GHz flux density can be converted to a total energy with no further assumptions about the burst spectral properties.

Figure 7 shows the FRB 121102 burst energy cumulative distribution as seen by the VLA and calculated from prior observations by Arecibo and the GBT (Spitler et al. 2016; Scholz et al. 2016). The latter two energy distributions are scaled from the fluence by assuming that all burst energy is included by the observation. This likely underestimates the burst energy for some bursts by a factor of a few, but should not affect the slope of the distribution.

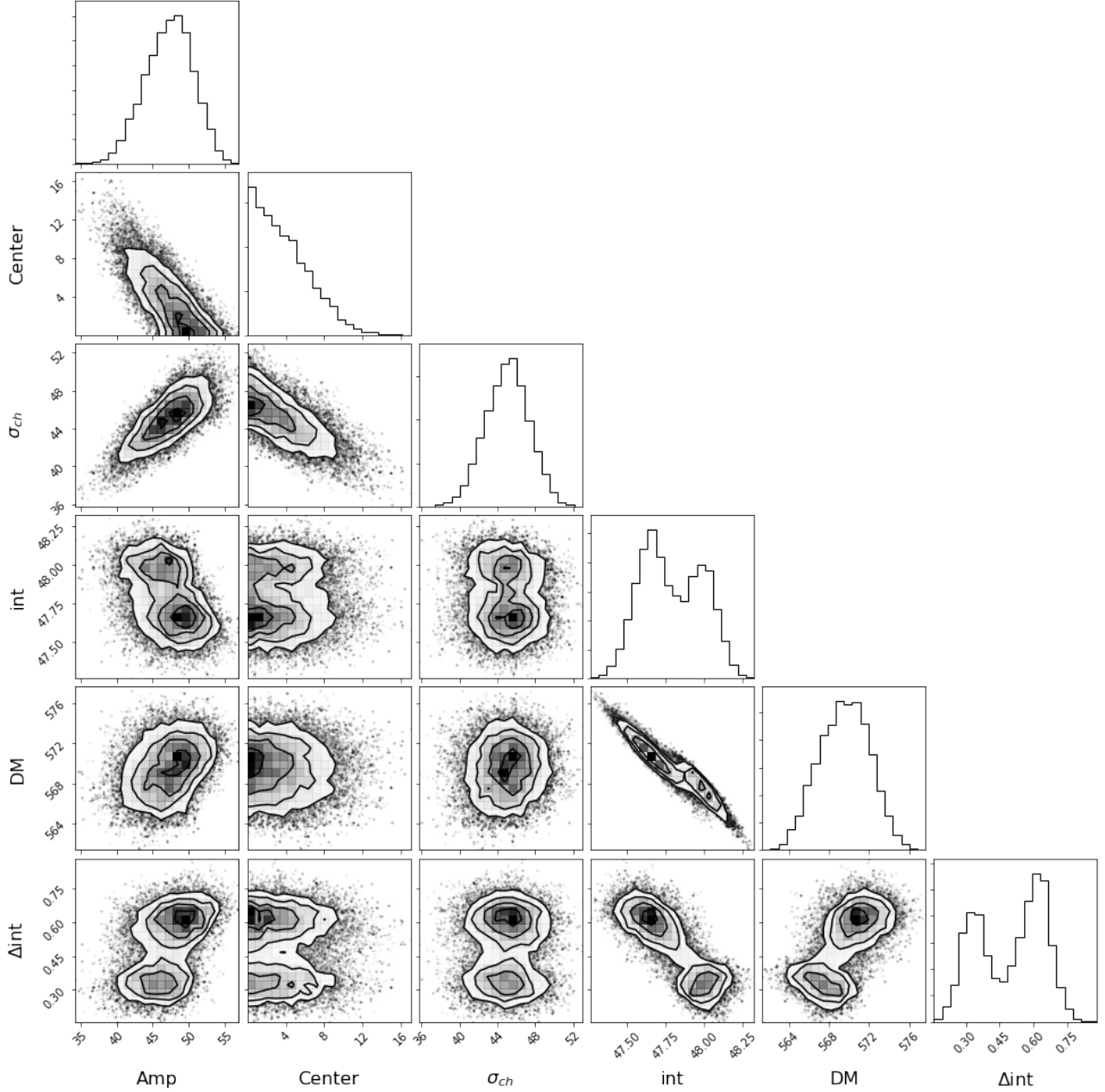
Broadly, the range of energies shown in Figure 7 is defined by the sensitivity of the telescope and the duration of the observing campaign. The VLA distribution shown represents the 27-hour, late-2016 campaign, which is longer than the single-dish campaigns presented here. However, since the VLA observations have a relatively coarse time resolution, it is not as sensitive as Arecibo and the GBT, so no bursts with energetics less than  $3 \times 10^{38}$  erg were detected. The rate upper limit (95% confidence) from the early-2016 VLA campaign shows that even identical observing campaigns have different detection rates.

We modeled the differential energy distribution,  $dN/d\log E$ , again using a Poisson detection probability with a rate function  $\lambda = AE^\alpha$ . Rather than trying to estimate a completeness limit for each energy distribution, we assume an effective detection limit of 0.9 times the weakest burst detected; the best-fit slope is weakly dependent on this, but general conclusions are robust. We directly sampled the likelihood distribution to estimate





**Figure 4.** Spectra and best-fit model for nine bursts seen by the VLA from 2.5 to 3.5 GHz in the peak, dedispersed 5 ms integration. Starting at the top left (moving right and then down), they correspond to bursts on MJDs 57623, 57633.68, 57633.70, 57638, 57643, 57645, 57646, 57648, and 57649. The solid line is a best-fit Gaussian model found through modeling.



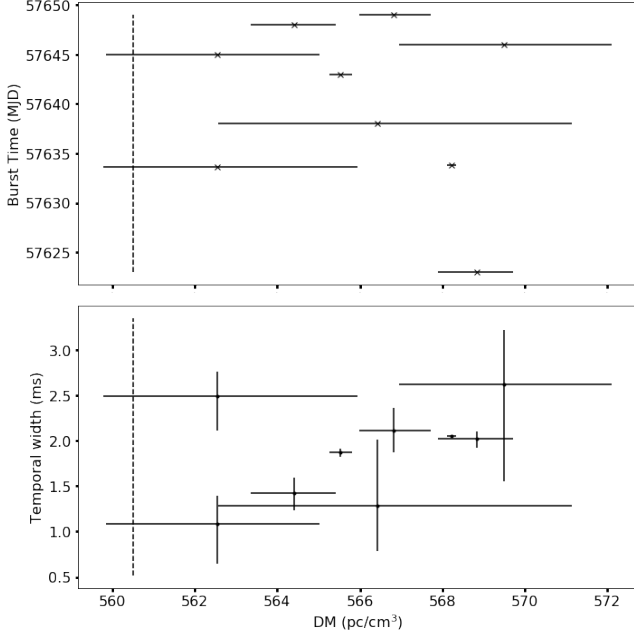
**Figure 5.** A scatterplot matrix that shows the correlation of every pair of parameters in the MCMC run for burst 57646. The equivalent plot for the other bursts show normally-distributed samples, but this burst shows significant structure in the samples. The parameters “Amp”, “Center”, and “ $\sigma_{ch}$ ” refer to the Gaussian spectral shape, while “DM”, “int”, and “W” refer to the temporal shape. The parameters are given in units of channels (4 MHz) and integrations (5 ms).

a best slope of  $\alpha_{\text{VLA}} = -0.7^{+0.3}_{-0.4}$ ,  $\alpha_{\text{Arecibo}} = -0.8^{+0.3}_{-0.5}$ , and  $\alpha_{\text{GBT}} = -0.8^{+0.4}_{-0.5}$  (68% confidence interval). Expressed as a power law function in  $dN/dE$ , the slope index is approximately  $-1.7$ .

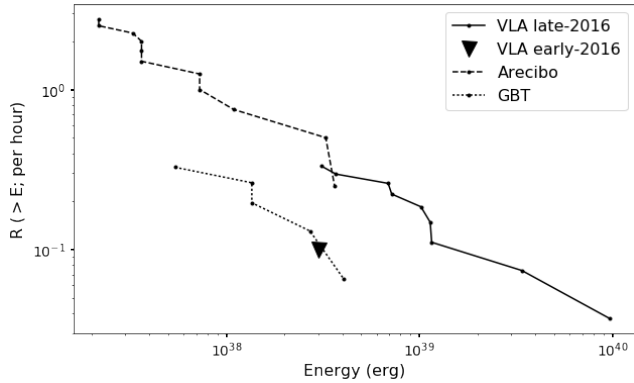
#### 4. DISCUSSION

##### 4.1. Burst Spectra

We present the first simultaneous detection of an FRB with multiple telescopes and over frequencies from 1.2 to 3.5 GHz. The flux density of the Arecibo burst detected on MJD 57648 is an order of magnitude less than that seen by the VLA. At the same time, three other bursts from FRB 121102 had similar observing coverage but were not detected simultaneously. This is consistent with the spectral structure observed within the 3 GHz



**Figure 6.** (Top) Burst time versus DM for nine 3 GHz bursts detected by the VLA. The DM scale for both panels is shown at the bottom. The 68% confidence interval on DM is shown with a bar and the dashed line shows the best-fit DM =  $560.5 \text{ pc cm}^{-3}$  inferred from Arecibo observations at 1.4 GHz (Hessels et al, in prep). (Bottom) Temporal width versus burst DM for the same nine VLA bursts.



**Figure 7.** The cumulative burst energy distributions for the VLA, Arecibo, and GBT bursts are shown with dots connected by solid, dashed, and dotted lines, respectively. The Arecibo distribution is derived from the 11 1.4 GHz bursts reported in Spitler et al. (2016) and the GBT distribution is derived from the 5 1.4 GHz bursts reported in Scholz et al. (2016). An upper limit from the VLA nondetection in early 2016 is shown as a triangle.

VLA band, which is typically limited to a Gaussian envelope with width of roughly 500 MHz. This confirms earlier results at 1.4 GHz (Spitler et al. 2016; Scholz

et al. 2016) with a wider bandwidth and extends the presence of this phenomenon to 3 GHz.

If FRB 121102 burst spectra are typical of the larger FRB population, then population models need to be generalized beyond the assumption of a spectral power law. Most obviously, FRB 121102 implies that future multi-telescope searches for bursts are unlikely to simultaneously detect bursts in different bands (c.f. Sallmen et al. 1999). Second, since burst spectra have limited bands, the burst detection rate at one frequency may not represent that at another. Rate estimates for the FRB population will need to explicitly account for the frequency-dependent rate. Finally, we note that the odds of detecting a burst will improve with bandwidth beyond the gain in sensitivity, since bandwidths wider than the FRB characteristic width are more likely to cover the burst envelope. However, FRB search algorithms will need to be modified to search for bursts with spectral width less than the full bandwidth.

After correcting for barycentric and typical dispersion delays, the VLA/Arecibo burst spectrum shows a residual temporal drift. This drift can be interpreted as an excess DM, but contemporaneous observations at 1.4 GHz measured a significantly lower DM ( $560.5 \text{ pc cm}^{-3}$ ; Hessels et al, in prep). This excess DM is comparable to the burst-to-burst variation in DM within the whole sample of VLA 3 GHz bursts. This variation had been noted in earlier 1.4 GHz Arecibo observations of FRB 121102 (Scholz et al. 2016), but the VLA bursts demonstrate that this burst-specific DM is seen up to 3 GHz.

One interpretation for the burst-dependent DM is that the bursts have a frequency-dependent pulse shape. Pulsar pulse shapes evolve with frequency on GHz frequency scales, presumably due to changes in the beam shape (Lyne & Manchester 1988). The burst-dependent DM changes are on the order of 1% of the total DM (equivalent to a delay rate up to  $2 \text{ ms GHz}^{-1}$ ), which is much larger than typically observed (Lentati et al. 2017). The systematic offset of the VLA 3 GHz DM from that observed by Arecibo suggests that any such burst frequency evolution must be larger at higher frequencies. Our modeling of the VLA burst dynamic spectra show a weak correlation between larger apparent DM and pulse width. This could be a signature of unresolved spectrotemporal structure in the VLA burst spectra that biases the measured DM.

The burst-specific DM structure could be intrinsic to the emission mechanism or induced during propagation. Radio waves can be modified in a variety of ways (e.g., scintillation, scattering) and the duration of the emission in the source frame is not known (Cordes et al. 2016).

Cordes et al. (2017) describe a model where plasma lensing near the source of an FRB can magnify radio emission by orders of magnitude. That model argues that the simultaneous appearance of a burst from 1.2 to 3.5 GHz requires a lens with the focal frequency be  $\gtrsim 3.5$  GHz. This gives a constraint on the combination

$$\frac{DM_0}{a_{AU}^2} \frac{d_{sl} d_{lo}/d_{so}}{1 \text{ kpc}} \gtrsim (3.5/39.1)^2$$

where  $DM_0$ ,  $a$ , and  $d_{sl, lo, so}$  are the electron column density of the lens, size of the lens, and lensing distances, respectively. A lens placed at  $d_{sl} \approx 1$  kpc is consistent with typical properties of the interstellar medium. Lenses that are influenced by the source would be much nearer to the source  $d \ll 1$  kpc, requiring large values of  $DM_0/a_{AU}^2$ .

#### 4.2. Energy Distribution

Our refined analysis shows that FRB 121102 has even higher isotropic energies than previously reported for FRB 121102, reaching energies  $E_{\text{max}} \approx 10^{40}$  erg. FRB 121102 has isotropic energies orders of magnitude higher than the strongest Crab giant pulses ( $\sim 10^{35}$  erg) and comparable brightness temperatures ( $T_{b, \text{Crab}} \sim 10^{41}$  K versus  $T_{b, \text{FRB121102}} \sim 10^{38}$  K; Hankins et al. 2003; Katz 2014).

In §3.4 we compare the energy distribution of the 9 VLA bursts to those reported earlier by Arecibo and the GBT. All three energy distributions can be characterized as a power law in  $dN/dE$  with slope  $\sim -1.7$ . This slope is seen even though the burst rate varies by almost an order of magnitude between observations and frequencies cover both 1.4 and 3 GHz bands. This suggests that the slope is related to the underlying physical process, rather than the burst detection rate at any given time. The value of the energy distribution slope is similar to that observed in magnetars (Göğüş et al. 2000; Scholz & Kaspi 2011). This similarity has been used to argue that bursts are driven by self-regulated critical phenomena (slope =  $-5/3$ ; Aschwanden 2011).

#### 4.3. Flux Distribution

While the surprising brightness of new FRB discoveries was once seen as a cause for concern (i.e., a sign of terrestrial interference), we may be able to use this to infer properties of the overall population (Lorimer et al. 2007; Ravi et al. 2016). When modeling the flux distribution as a power law, an isotropic distribution of sources will have a power law index,  $\alpha$ , of  $-1.5$ . Prior to the measurement of the first GRB redshift, deviations from this Euclidean distribution were used to infer their cosmological distribution (e.g., the  $V/V_{\text{max}}$  test; Mao

& Paczynski 1992; Fenimore & Bloom 1995). Multiple studies have suggested that FRBs have a sub-Euclidean flux distribution ( $-0.5 < \alpha < -0.9$ ; Vedantham et al. 2016; Li et al. 2016; Lawrence et al. 2016). Others have noted that the flux distribution is not well modeled as a power law, such that  $\alpha$  is effectively sensitivity (and telescope) dependent (Oppermann et al. 2016; CHIME Scientific Collaboration et al. 2017).

We now know that FRB 121102 is detectable with the VLA out to  $z = 0.7$ , which suggests that cosmological effects are even more important for understanding properties of the FRB population. CHIME Scientific Collaboration et al. (2017) demonstrate how redshift space distortions, time dilation, and spectral index (“k-correction”) effects can flatten the flux distribution. To this complex scene, we add the fact that FRB 121102 burst spectra are not defined by a spectral index, which suggests that k-corrections are likely to be very difficult to calculate in practice. If the FRB burst rate is higher at low (rest) radio frequencies, then high-redshift FRBs will have an lower (redshifted) rate that flattens the  $dN/dF$  distribution in a manner similar to a k-correction. However, propagation effects are expected to suppress the FRB rate at sub-GHz frequencies (Rajwade & Lorimer 2017; Chawla et al. 2017) as well. The VLA and Arecibo observed intensively in a coordinated fashion during the late-2016 campaign and a comparison of their relative rates will be presented elsewhere.

#### 4.4. Repetition

The chance of detecting a burst from FRB 121102 with the VLA changes substantially on day-month timescales. A major outstanding question is whether this time-variable rate is driven by an intrinsic process (Katz 2016) or extrinsic (propagation) effects in the host galaxy or intergalactic medium (Cordes et al. 2017). However, the temporal distribution of VLA detections alone shows that the FRB 121102 burst rate is significantly correlated on short timescales. That is, detecting a burst predicts that you are more likely to find another burst soon thereafter and not detecting a burst predicts you will likely not detect another burst soon thereafter.

The burst time distribution is consistent with a “red spectrum” (Connor et al. 2016a), although the observations presented here are only sensitive to timescales of days (time between observations) and months (time between campaigns). As has been noted elsewhere, this statistical property would weaken previous constraints on repetition (Petroff et al. 2015; Law et al. 2015). If this statistical property describes other FRBs, then wide, shallow surveys are the preferred strategy for blind de-

tection of FRBs. Also, since “bursts predict bursts”, many short observations are more likely to detect an FRB than a single long observation of the same total length and all new FRB detections should be immediately and intensively followed up.

#### 4.5. Volumetric Rate of FRB Sources

Thornton et al. (2013) estimated an isotropic rate of FRB events of  $10^{-3} \text{ galaxy}^{-1} \text{ yr}^{-1}$  by calculating the number of Milky Way-type galaxies out to a distance implied by assuming all extragalactic DM originates in the IGM ( $z \approx 0.9$  for  $\text{DM} \approx z \times 900 \text{ pc cm}^{-3}$ ; Ioka 2003; Inoue 2004). This rate is high (comparable to the rate of core-collapse supernovae; Diehl et al. 2006), which has been used to argue that FRBs are likely not associated with other classes of transient, such as LGRBs or rare subclasses of SNe (Woosley & Bloom 2006). However, since this rate depends on telescope sensitivity, it is not appropriate to compare to other estimates based on source counts.

Here, we reevaluate the FRB volumetric rate by assuming FRB 121102 is a prototype of the class and recasting it as a volumetric rate of FRB sources (i.e., a birth rate). Crudely, the volumetric rate of FRB sources can be defined as  $R_{\text{FRB}} = R_p / (N_r \Omega_b V(z))$ , where  $R_p$  is the projected FRB rate,  $N_r$  is the number of bursts per source in a typical lifetime,  $\Omega_b$  is the beaming fraction, and  $V(z)$  is the comoving volume out to redshift  $z$ . The newest estimates of the projected FRB rate are  $2 \times 10^3 \text{ sky}^{-1} \text{ day}^{-1}$  at high Galactic latitudes and flux densities brighter than 1 Jy ms (Lawrence et al. 2016; Champion et al. 2016; Rane et al. 2016). There is relatively little constraint on the beaming factor, but Galactic pulsars have beaming solid angle fractions on the order of 10% (Tauris & Manchester 1998).

FRB 121102 has an absolute energy scale that can be used to estimate a horizon scale for the typical FRB survey sensitivity. At the typical Parkes survey parameters (e.g., Champion et al. 2016), a survey with a fluence limit of 1 Jy ms can detect FRB 121102 at  $z \approx 1$ . The distance inferred from the largest DMs observed ( $\sim 1500 \text{ pc cm}^{-3}$ ; Champion et al. 2016) is somewhat higher, but that likely overestimates distance since DM is expected to have significant contributions from the host galaxy and intervening galaxies (Tendulkar et al. 2017; McQuinn 2014).

Using an assumed horizon redshift of 1, we estimate a rate of  $R_{\text{FRB}} \approx 5 \times 10^{-5} N_r^{-1} (0.1/\Omega_b) \text{ Mpc}^{-3} \text{ yr}^{-1}$ . This calculation ignores time dilation, which underestimates the burst rate by a factor of 2 at the redshift horizon. The narrow burst spectral structure seen for FRB 121102 also suggests that the FRB rate is under-

stated by a factor of a few by ignoring bursts that fall outside the observing band. Finally, we note that this estimate assumes that Galactic latitude and propagation effects do not significantly affect the detectability of FRBs, although scintillation and scattering likely play significant roles (Macquart & Johnston 2015; Cordes et al. 2017).

Despite its many assumptions, this volumetric rate is more appropriate for comparison to rates used for other classes of transient. The LGRB and SLSN-I rates are  $10^{-7}$  and  $10^{-8} \text{ Mpc}^{-3} \text{ yr}^{-1}$ , respectively (Guetta & Della Valle 2007; Gal-Yam 2012). Comparing the FRB rate to these rates implies that FRB-emitting sources can be generated in LGRBs and SLSN-I if FRBs repeat  $5 \times 10^2$  and  $5 \times 10^3$  times<sup>4</sup>, respectively (assuming burst energies of  $\sim 10^{40} \text{ erg}$ ). Metzger et al. (2017) use the young magnetar model for FRBs to calculate the maximum number of bursts that can be powered its magnetic field as  $\approx 300$ . Assuming this model for FRBs, LGRBs, and SLSN-I, the rate comparison suggest that the LGRB rate is most consistent with the FRB rate, although there is significant room for adjustment in the models and rate estimate.

## 5. CONCLUSIONS

The recent precision localization of FRB 121102 has helped identify a host galaxy, measure its distance, and establish it as a member of a truly new class of astrophysical source. With its cosmological distance firmly established, FRB 121102 now serves as a new kind of standard by which FRBs are defined. That fact, combined with a rich data set of many bursts, allows us to apply this source to more general questions about the FRB population.

We presented the first multi-telescope detection (Arecibo and VLA) of an FRB. By detecting this burst from 1.2 to 3.5 GHz, we have demonstrated that some bursts have broad spectral structure. However, the energy of that burst was dominated by the VLA 3 GHz observing band and three other VLA bursts are undetected by Arecibo. This demonstrates the burst spectra are poorly described by a power law with a single spectral index. Independently, we modelled dynamic spectra within the VLA 3 GHz band and show that most bursts can be characterized by a Gaussian envelope of width  $\sim 500 \text{ MHz}$ . This modeling also shows that the apparent DM changes from burst to burst and that DM is biased above the long-term average measured at 1.4 GHz. The nature of this DM change is not known, but could be ex-

<sup>4</sup> Nicholl et al. (2017) provide a similar estimate of FRB birth rate and reach a similar conclusion.

plained by strong frequency-dependent profile evolution or unresolved spectrottemporal in the bursts.

With a characteristic burst spectrum, we can estimate total (apparent) radio energy per burst. The cumulative energy distribution is characterized by a power law in  $dN/dE$  with slope of  $\sim -1.7$ . The amplitude of this power law changes significantly between observational campaigns. The stochastic nature of the burst rate suggests that past constraints on FRB repetition are weaker than previously inferred. The relatively narrow spectral structure, flat energy distribution, and variable burst rate of FRB 121102 suggests that repeated observations, wide bandwidth, and large instantaneous field of view all improve the odds of FRB discovery.

Assuming that FRB 121102 is representative of the FRB population, we calculate a volumetric birth rate of FRB sources that does not depend on telescope sensitivity. We estimate a volumetric rate of FRB sources of  $R_{\text{FRB}} \approx 5 \times 10^{-5} (0.1/\Omega_b) \text{ Mpc}^{-3} \text{ yr}^{-1}$ , which is appropriate for FRB 121102-like burst energies of  $10^{40}$  erg that are detectable out to  $z = 1$ . This rate is broadly consistent with models of FRBs from young pulsars or magnetars born in superluminous supernovae or long gamma-ray bursts, if the typical FRB repeats on the order of  $10^3$  times over its lifetime.

New, arcsecond-scale localizations will be critical to refining the picture presented here and constraining models of FRB origin. FRB 121102 was localized within hours by a prototype version of *realfast* and an expanded *realfast* system is now under construction. This platform will search a TB/hour data stream in real time in parallel with ongoing VLA observations, potentially detecting and localizing multiple FRBs per year.

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*Software:* *rtpipe*

#### REFERENCES

- Aschwendan, M. J. 2011, *SoPh*, 274, 99
- Beloborodov, A. M. 2017, ArXiv e-prints, [arXiv:1702.08644 \[astro-ph.HE\]](#)
- Champion, D. J., Petroff, E., Kramer, M., et al. 2016, *MNRAS*, 460, L30
- Chatterjee, S., Law, C. J., Wharton, R. S., et al. 2017, *Nature*, 541, 58
- Chawla, P., Kaspi, V. M., Josephy, A., et al. 2017, ArXiv e-prints, [arXiv:1701.07457 \[astro-ph.HE\]](#)
- CHIME Scientific Collaboration, Amiri, M., Bandura, K., et al. 2017, ArXiv e-prints, [arXiv:1702.08040 \[astro-ph.HE\]](#)
- Connor, L., Pen, U.-L., & Oppermann, N. 2016a, *MNRAS*, 458, L89
- Connor, L., Sievers, J., & Pen, U.-L. 2016b, *MNRAS*, 458, L19
- Cordes, J. M., & Lazio, T. J. W. 2002, ArXiv Astrophysics e-prints, [arXiv:astro-ph/0207156](#)
- Cordes, J. M., & McLaughlin, M. A. 2003, *ApJ*, 596, 1142
- Cordes, J. M., & Wasserman, I. 2016, *MNRAS*, 457, 232
- Cordes, J. M., Wharton, R. S., Spitler, L. G., Chatterjee, S., & Wasserman, I. 2016, ArXiv e-prints, [arXiv:1605.05890 \[astro-ph.HE\]](#)
- Cordes, J. M., Freire, P. C. C., Lorimer, D. R., et al. 2006, *ApJ*, 637, 446
- Cordes et al. 2017, in prep
- Diehl, R., Halluin, H., Kretschmer, K., et al. 2006, *Nature*, 439, 45

- Dokuchaev, V. I., & Eroshenko, Y. N. 2017, ArXiv e-prints, 1701.02492, [arXiv:1701.02492 \[astro-ph.HE\]](#)
- Fenimore, E. E., & Bloom, J. S. 1995, *ApJ*, 453, 25
- Foreman-Mackey, D. 2016, *The Journal of Open Source Software*, 24
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, 125, 306
- Fuller, J., & Ott, C. D. 2015, *MNRAS*, 450, L71
- Gal-Yam, A. 2012, *Science*, 337, 927
- Goodman, J., & Weare, J. 2010, *Communications in applied mathematics and computational science*, 5, 65
- Gögüş, E., Woods, P. M., Kouveliotou, C., et al. 2000, *ApJL*, 532, L121
- Guetta, D., & Della Valle, M. 2007, *ApJL*, 657, L73
- Hankins, T. H., Kern, J. S., Weatherall, J. C., & Eilek, J. A. 2003, *Nature*, 422, 141
- Inoue, S. 2004, *MNRAS*, 348, 999
- Ioka, K. 2003, *ApJL*, 598, L79
- Kashiyama, K., & Murase, K. 2017, ArXiv e-prints, 1701.04815, [arXiv:1701.04815 \[astro-ph.HE\]](#)
- Katz, J. I. 2014, *PhRvD*, 89, 103009
- . 2016, *ApJ*, 826, 226
- Kulkarni, S. R., Ofek, E. O., Neill, J. D., Zheng, Z., & Juric, M. 2014, *ApJ*, 797, 70
- Law, C. J., Bower, G. C., Burke-Spolaor, S., et al. 2017, in *American Astronomical Society Meeting Abstracts*, Vol. 229, *American Astronomical Society Meeting Abstracts*, 330.02
- Law, C. J., Bower, G. C., Burke-Spolaor, S., et al. 2015, *ApJ*, 807, 16
- Lawrence, E., Vander Wiel, S., Law, C. J., Burke Spolaor, S., & Bower, G. C. 2016, ArXiv e-prints, [arXiv:1611.00458 \[astro-ph.HE\]](#)
- Lentati, L., Kerr, M., Dai, S., et al. 2017, *MNRAS*, 466, 3706
- Li, L., Huang, Y., Zhang, Z., Li, D., & Li, B. 2016, ArXiv e-prints, [arXiv:1602.06099 \[astro-ph.HE\]](#)
- Lorimer, D. R., Bailes, M., McLaughlin, M. A., Narkevic, D. J., & Crawford, F. 2007, *Science*, 318, 777
- Lunnan, R., Chornock, R., Berger, E., et al. 2014, *ApJ*, 787, 138
- Lyne, A. G., & Manchester, R. N. 1988, *MNRAS*, 234, 477
- Macquart, J.-P., & Johnston, S. 2015, *MNRAS*, 451, 3278
- Mao, S., & Paczynski, B. 1992, *ApJL*, 388, L45
- Marcote, B., Paragi, Z., Hessels, J. W. T., et al. 2017, *ApJL*, 834, L8
- McQuinn, M. 2014, *ApJL*, 780, L33
- Metzger, B. D., Berger, E., & Margalit, B. 2017, ArXiv e-prints, 1701.02370, [arXiv:1701.02370 \[astro-ph.HE\]](#)
- Modjaz, M., Kewley, L., Kirshner, R. P., et al. 2008, *AJ*, 135, 1136
- Nicholl, M., Williams, P. K. G., Berger, E., et al. 2017, ArXiv e-prints, [arXiv:1704.00022 \[astro-ph.HE\]](#)
- Oppermann, N., Connor, L. D., & Pen, U.-L. 2016, *MNRAS*, 461, 984
- Perrott, Y. C., Scaife, A. M. M., Green, D. A., et al. 2013, *MNRAS*, 429, 3330
- Petroff, E., Johnston, S., Keane, E. F., et al. 2015, *MNRAS*, 454, 457
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, 594, A13
- Popov, S. B., & Pshirkov, M. S. 2016, *MNRAS*, 462, L16
- Rajwade, K. M., & Lorimer, D. R. 2017, *MNRAS*, 465, 2286
- Rane, A., Lorimer, D. R., Bates, S. D., et al. 2016, *MNRAS*, 455, 2207
- Ransom, S. M. 2001, PhD thesis, Harvard University
- Ravi, V., Shannon, R. M., Bailes, M., et al. 2016, ArXiv e-prints, [arXiv:1611.05758 \[astro-ph.HE\]](#)
- Sallmen, S., Backer, D. C., Hankins, T. H., Moffett, D., & Lundgren, S. 1999, *ApJ*, 517, 460
- Scholz, P., & Kaspi, V. M. 2011, *ApJ*, 739, 94
- Scholz, P., Spitler, L. G., Hessels, J. W. T., et al. 2016, ArXiv e-prints, [arXiv:1603.08880 \[astro-ph.HE\]](#)
- Spitler, L. G., Cordes, J. M., Hessels, J. W. T., et al. 2014, *ApJ*, 790, 101
- Spitler, L. G., Scholz, P., Hessels, J. W. T., et al. 2016, *Nature*, 531, 202
- Tauris, T. M., & Manchester, R. N. 1998, *MNRAS*, 298, 625
- Tendulkar, S. P., Bassa, C. G., Cordes, J. M., et al. 2017, *ApJL*, 834, L7
- Thompson, C. 2017, ArXiv e-prints, [arXiv:1703.00393 \[astro-ph.HE\]](#)
- Thornton, D., Stappers, B., Bailes, M., et al. 2013, *Science*, 341, 53
- Vedantham, H. K., Ravi, V., Hallinan, G., & Shannon, R. M. 2016, *ApJ*, 830, 75
- Woosley, S. E., & Bloom, J. S. 2006, *ARA&A*, 44, 507
- Zhang, B. 2017, ArXiv e-prints, 1701.04094, [arXiv:1701.04094 \[astro-ph.HE\]](#)