

Research Summary

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I. INTRODUCTION TO COSMIC MICROWAVE BACKGROUND(CMB)

Anisotropies in the cosmic mircrowave background(CMB) radiation have provided a weath of information about cosmological model that describes the contents and evolution of the universe[1]. We extract multipoles from CMB temperature anisotropies(T) and polarization anisotropies(E,B) which traces the behaviour at 400 000 years of cosmic time, as well as from the signals imprinted at later times due to scattering from galaxy clusters, from the motion of electrons ionized, and from gravitationl lensing. FIG 1. shows the full sky temperature anisotropies from Planck.

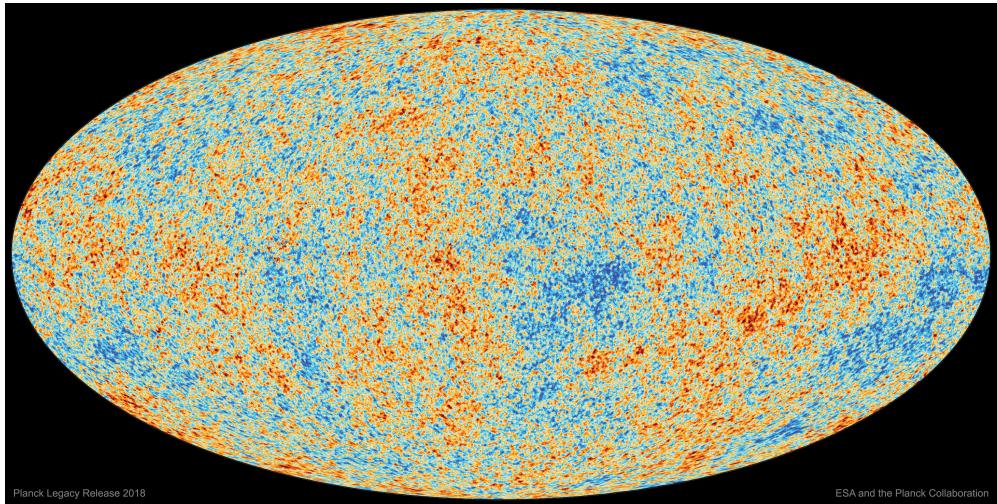


FIG. 1: The 2018 Planck map of the temeprature anisotropies of the CMB

II. CMB LENSING RECONSTRUCTION ON THE CURVED SKY

Cosmic Microwave Background(CMB) photons are gravitationlly lensed by the large scale mass distribution. This effect is called CMB lensing. CMB temperature anisotropies(T) and polarization anisotropies(E,B) are distorted by CMB lensing. We can use deflection field or lensing potential to discribe CMB lensing[2]. It is expected to amongst the most powerful

cosmological tools for ongoing and upcoming CMB experiments which will help find out primordial gravitational waves, constrain cosmological parameters, constrain neutrino mass and understand the properties of dark energy and . So it is important to reconstruct the CMB lensing according to the observed CMB temperature and polarization anisotropies. The process of reconstructing lensing potential is called CMB lensing reconstruction.

There are two aspects where we can see CMB lensing effects: 1.CMB lensing modulates CMB temperature and polarization power spectra(2 point power spectrum). Statistical anisotropy is induced. 2.CMB lensing produces higher-order correlations between the multipole moments. Off-diagonal mode-coupling between map harmonics can be seen. The off-diagonal mode coupling is proportional to lensing potential.[3]

$$\left\langle a_l^m b_{l'}^{m'} \right\rangle = C_l^{ab} \delta_{ll'} \delta_{m-m'} (-1)^m \quad (1)$$

$$\left. \left\langle a_l^m b_{l'}^{m'} \right\rangle \right|_{\text{lens}} = C_l^{ab} \delta_{ll'} \delta_{m-m'} (-1)^m + \sum_{LM} (-1)^M \begin{pmatrix} l & l' & L \\ m & m' & -M \end{pmatrix} f_{lLl'}^\alpha \phi_L^M \quad (2)$$

Eq(1) and Eq(2) are the covariance of unlensed and lensed CMB multipoles respectively. a_l^m and b_l^m are multipoles for T, E or B-mode multipoles. ϕ_l^m is lensing potential multipole. As we can see, the multipole covariance gained an extra term which is proportional to lensing potential and Wigner-3j symbols.

For primordial gravitational waves, CMB lensing can be an important source of confusion[2]. Density perturbations in the linear regime generate only the so-called E-mode polarization[4]. Lensing converts E-mode polarization to B-mode polarization.[5]. Gravitational waves can also produce B-mode polarization[6]. Removing the lensing-generated B-mode is called delensing.

To obtain lensing potential, we can take quadratic combinations of CMB fields and construct quadratic estimators with minimum variance[6]. On small angular scales $L > 10^2$, we are able to neglect the curvature of the sky and apply flat sky estimators[3]. Since lensing is most sensitive to the projected potential at $L < 10^2$ or several degrees on the sky, we need consider the curvature of the sky. In this study, we apply estimators for lensing potential on the curved sky provided by Ref[7] as shown is Eq(3):

$$d_L^{\alpha M} = \frac{A_L^\alpha}{\sqrt{L(L+1)}} \sum_{l_1 m_1} \sum_{l_2 m_2} (-1)^M \begin{pmatrix} l_1 & l_2 & L \\ m_1 & m_2 & -M \end{pmatrix} g_{l_1 l_2}^\alpha(L) a_{l_1}^{m_1} b_{l_2}^{m_2} \quad (3)$$

We reconstruct lensing potential and get noise properties for the full sky. This work is meant for improving the existing lensing reconstruction pipeline in Advanced Atacama Cosmology Telescope(AdvACT) and Simons Observatory(SO). This work now only works for the full sky using TT and EB estimators. And its performance is still limited by now, expected to be improved.

I've learned lensing reconstruction techniques from this work which are applied in the work related lensing reconstruction bias shown in the section 2 below.

III. BIAS TO CMB LENSING RECONSTRUCTION FROM TEMPERATURE ANISOTROPIES DUE TO REIONIZATION kSZ

In this work, we are investigating a bias to CMB lensing reconstruction from temperature anisotropies due to the reionization kSZ effect(kinematic Sunyaev-Zel'dovich effect) based on simulation.

There are several ongoing and upcoming experiments, including Advanced Atacama Cosmology Telescope(AdvACT)[8], the South Pole Telescope-3G(SPT-3G)[9], the Simons Observatory[10], and CMB Stage-4(CMB-S4)[11]. For these experiments, the CMB lensing power spectrum will be measured with signal-to-noise(S/N) > 100 . At this precision level, we are required to consider more about biases in CMB reconstruction. Most of the biases can be removed from primary CMB by a multifrequency component separation methods, but they don't work for kSZ effect, since kSZ effect preserves blackbody spectrum of the CMB.[12]

kSZ effect is the Doppler shift of CMB photons induced by Compton-scattering off moving electrons(bulk velocity).[13]. The kSZ signal has its own intrinsic non-Gaussianity and its correlation with CMB lensing field is non-zero.[14].

This kSZ effect produces a CMB temperature change, $\Theta^{\text{kSZ}}(\hat{\mathbf{n}}) = \Delta T^{\text{kSZ}}(\hat{\mathbf{n}})/T_{\text{CMB}}$,

$$\Theta^{\text{kSZ}}(\hat{\mathbf{n}}) = -\sigma \int \frac{d\eta}{1+z} e^{-\tau} n_e(\hat{\mathbf{n}}, \eta) \mathbf{v}_e \cdot \hat{\mathbf{n}} \quad (4)$$

[13],

where σ is the Compton scattering cross-section, τ is the optical depth to Compton scattering, n_e is the physical free number density, \mathbf{v}_e is the peculiar velocity of the electrons. kSZ anisotropies can be produced when large fluctuation in electron density ap-

pears. There are two epochs when they can be produced: 1.a “late-time” contribution from redshifts $0 < z < 6$ in which inhomogeneities are large due to gravitational growth of structure 2.earlier during the epoch of reionization(from first stars, $6 < z < 20$) when hydrogen gets ionized again by the ultraviolet radiation of the first structures and the fluctuations electron density are caused by the fluctuations in the ionization fraction.[13] [15]. It is expected to be correlated with lensing field.

Simone Ferraro and J. Colin Hill have investigated the case of “late time” kSZ in [13]. Accordin to their results, the bias induced by “late time” kSZ to CMB lensing auto-power spectrum measurements can be as large as approximately %1, %6, and %8 for Plank, CMB-S3, and CMB-S4, respectively, when using $l_{max} = 4000$, and about half of that for $l_{maxy} = 3000$. Thus, for CMB-S3 and CMB-S4 lensing measurements, the kSZ-induced bias cannot be neglected.

For the case of reionization kSZ, it could also contribute at some level. We are trying to estimate reionization-induced bias to CMB lensing reconstruction from temperature anisotropies based on lensed temperature anisotropies[16] and reionization kSZ map from Ref[15]. In this study, we apply flat-sky reconstruction pipeline to cutout patchy lensed temperature maps with reionization kSZ and without reionization kSZ given different noise levels, beam sizes and l_{max} . We compare their reconstruction lensing auto-powerspectra and see how much bias the reionization kSZ induces. I am now simulating the bias and trying to understand the bias quantitatively. Since I am now fixing the bugs in this pipeline, the results shown below is not the conclusion. This work is also expected to be helpful for seperating reionization kSZ signals from lensing signals.

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