A Search for the Neutrinoless Double Beta Decay of Xenon-136 with Improved Sensitivity from Denoising

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April 3, 2014

Outline

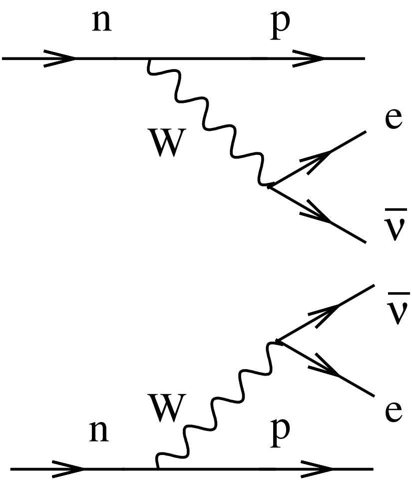
$\beta\beta 2\nu$ and $\beta\beta 0\nu$ Decay

The EXO-200 Detector

Denoising

Results

What is Double-Beta Decay?



Avignone et al., RMP 2008.

Feynman diagram for $\beta\beta 2\nu$ decay. Equivalent to two single-β decays:

$$2\mathrm{n} \rightarrow 2\mathrm{p} + 2\mathrm{e}^- + 2\bar{\nu}_\mathrm{e}$$

eg. $^{136}\mathrm{Xe} \rightarrow {}^{136}\mathrm{Ba} + 2\mathrm{e}^- + 2\bar{\nu}_\mathrm{e}$

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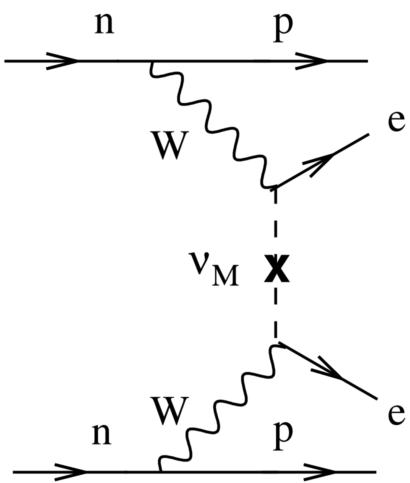

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$$\text{eg. } {}^{136}\mathrm{Xe} \to {}^{136}\mathrm{Ba} + 2\mathrm{e}^- + 2\bar{\nu}_\mathrm{e}$$

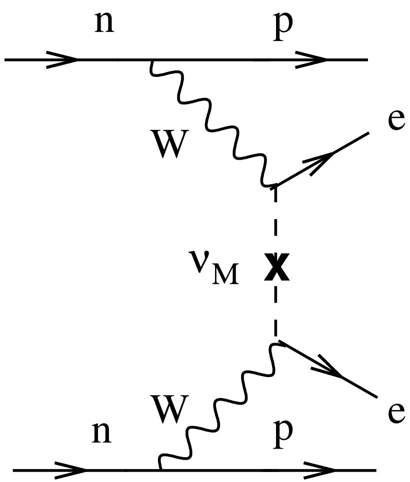
Feynman diagram for $\beta\beta 0\nu$ decay. Neutrinos annihilate each other:

$$2\mathrm{n} \to 2\mathrm{p} + 2\mathrm{e}^-$$

$$\text{eg. } {}^{136}\mathrm{Xe} \to {}^{136}\mathrm{Ba} + 2\mathrm{e}^-$$

$\beta\beta 2\nu$ is allowed in the Standard Model; $\beta\beta 0\nu$ is not.

Implications of Double-Beta Decay



Avignone et al., RMP 2008.

- Lepton number changes:

$$\Delta L = +2$$

- Neutrinos can convert to their own antiparticle:

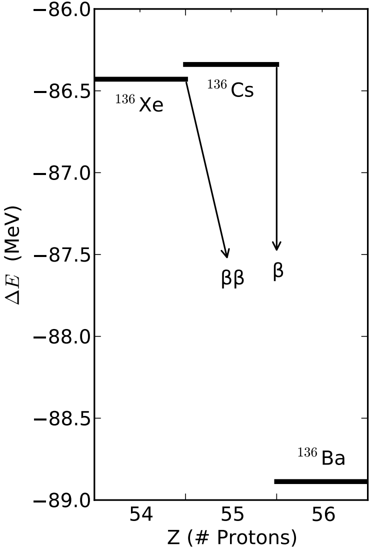
$$\bar{\nu}_R \rightarrow \nu_L$$

- Neutrinos have mass through a Majorana interaction:

$$-\frac{m_L}{2}\left(\overline{\Psi^c_L}\Psi_L + \overline{\Psi_L}\Psi^c_L\right)$$

$$-\frac{m_R}{2}\left(\overline{\Psi^c_R}\Psi_R + \overline{\Psi_R}\Psi^c_R\right)$$

The A = 136 Isobar



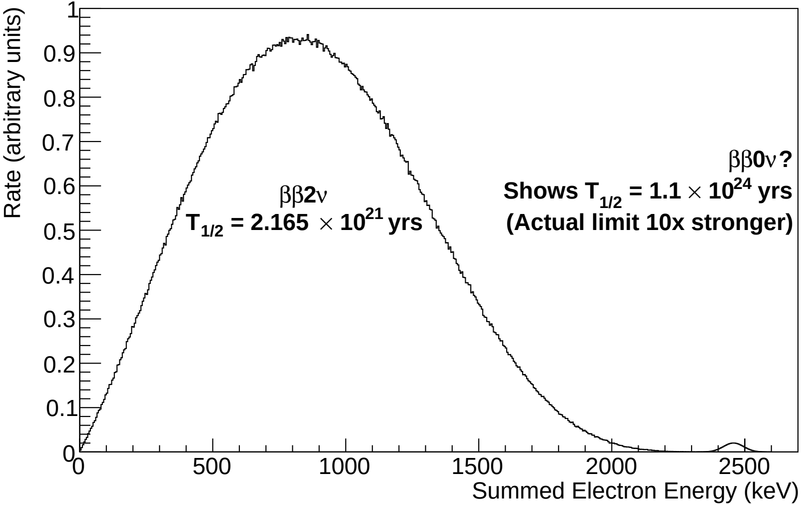
^{136}Cs undergoes single-β decay.

^{136}Xe cannot, due to energy conservation – but it can $\beta\beta$ decay through ^{136}Cs to ^{136}Ba.

The Q-value of $^{136}\mathrm{Xe} \rightarrow ^{136}\mathrm{Ba}$ is 2457.83 ± 0.37 keV, shared between all final products of the decay.

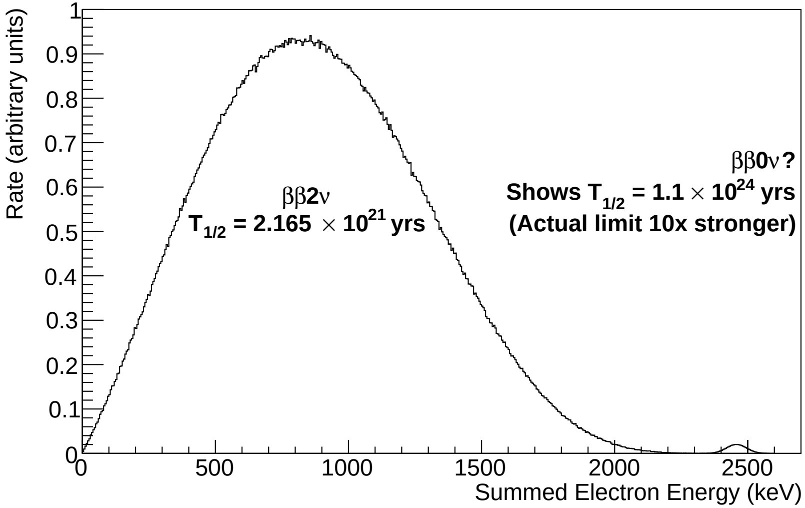
We observe energy in electrons; energy in neutrinos is lost.

Ideal Double-Beta Energy Spectrum



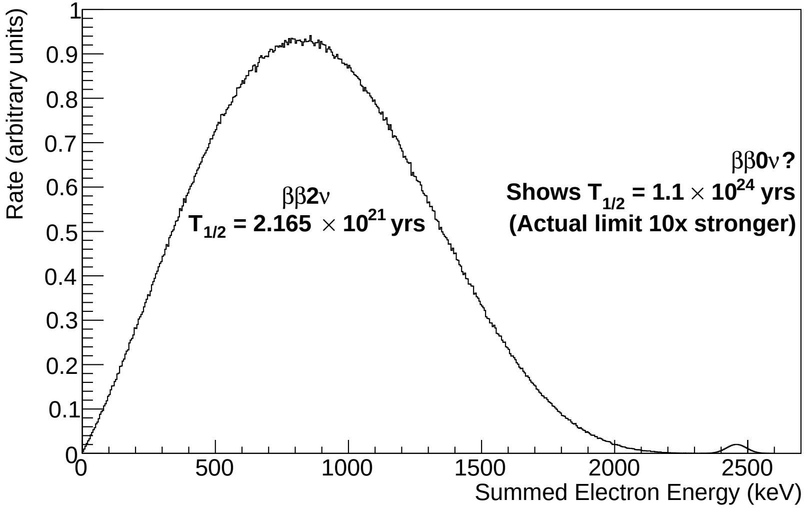
^{136}Xe $\beta\beta2\nu$ produces a smooth energy spectrum; "missing" energy carried off by neutrinos.

Ideal Double-Beta Energy Spectrum



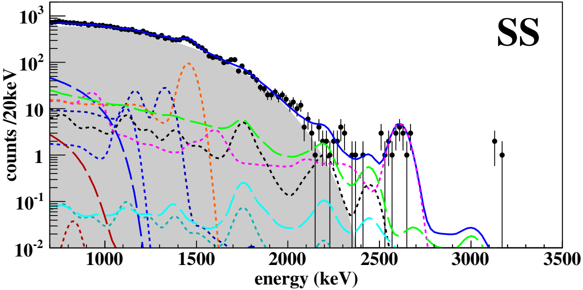
^{136}Xe $\beta\beta0\nu$ has no neutrinos, so no "missing" energy; mono-energetic peak at Q = 2458 keV.

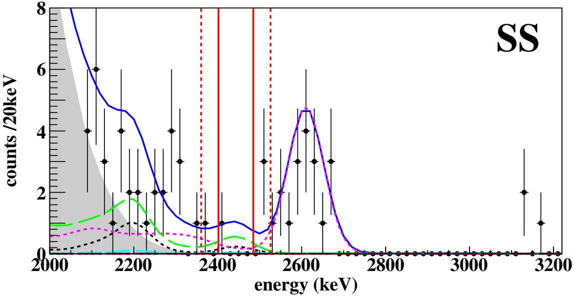
Ideal Double-Beta Energy Spectrum



If the $\beta\beta 0\nu$ peak exists, neutrinos have Majorana mass; peak height gives a measurement of that mass.

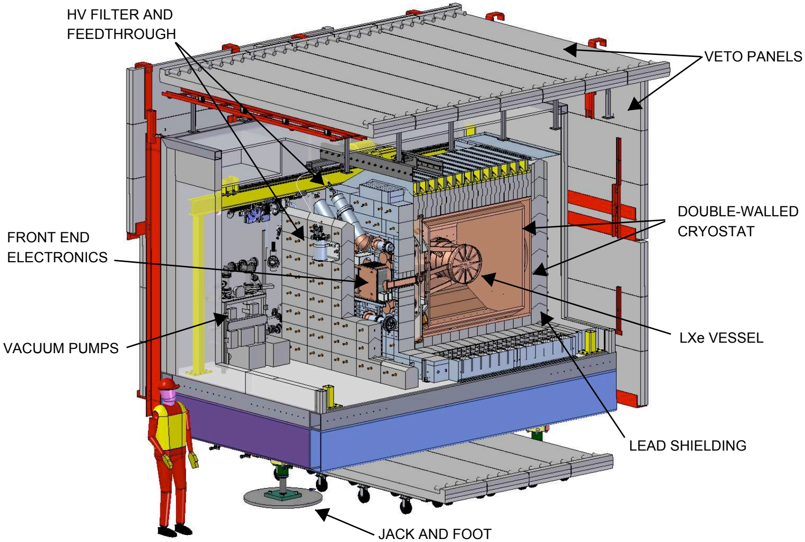
Observed Energy Spectrum (May 2012)





32.5 kg-yrs (vs. 99.8 kg-yrs for the present analysis)

The EXO-200 Detector



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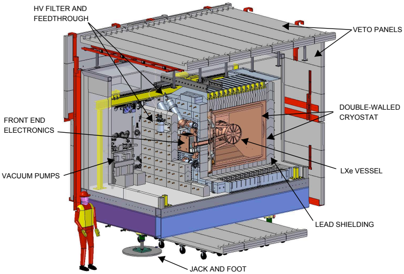
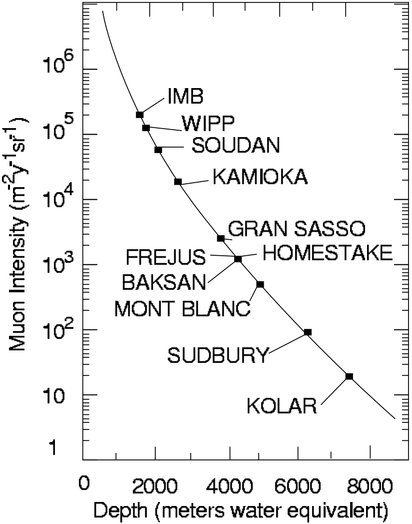


To search for rare decays, low background is key:

- Clean (low-radioactivity) materials surrounding TPC.
- Deep underground to avoid cosmogenics.

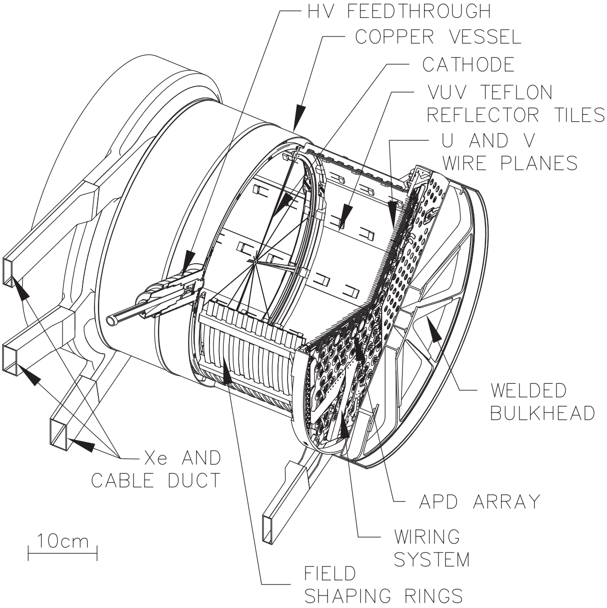


Esch et al., NIM A 2005.

EXO-200 TPC

110 kg of liquid xenon in active volume, enriched to 80.6% in ^{136}Xe, contained in a time projection chamber (TPC).

Xenon continuously circulates through purifiers outside of the cryostat.



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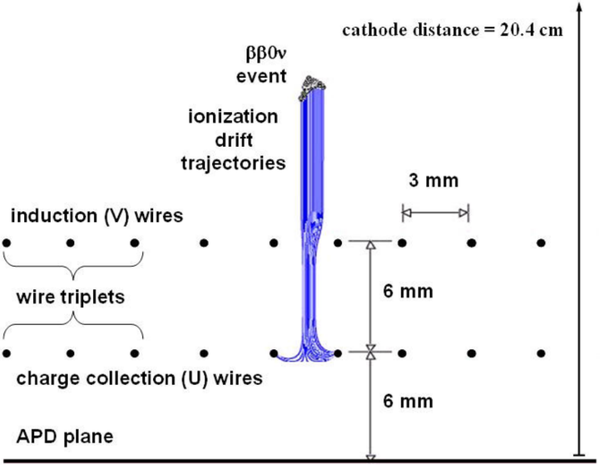
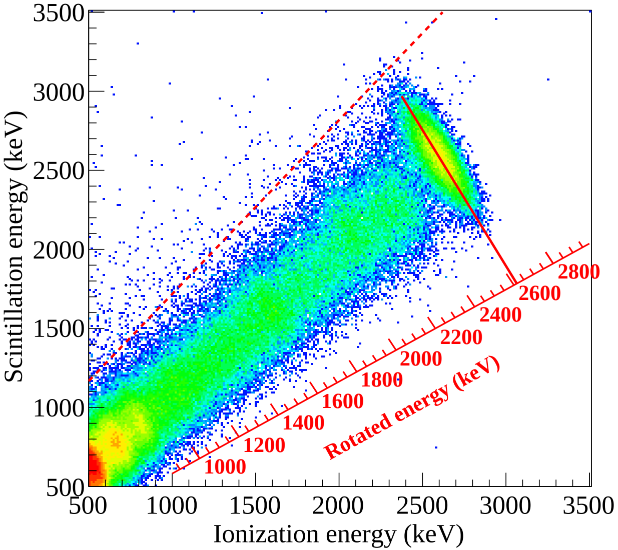


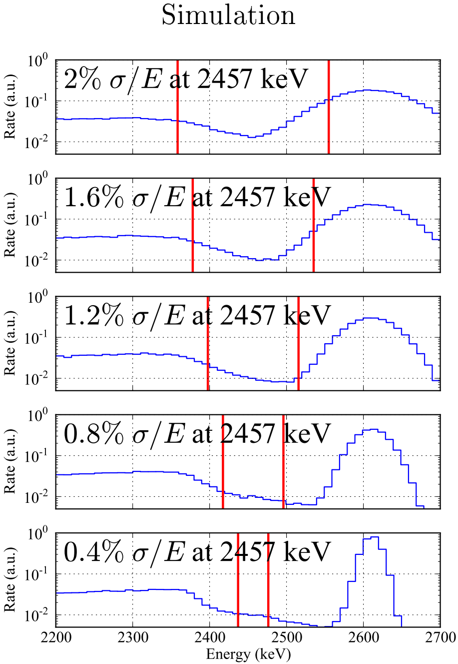
Charge drifts under an electric field and is collected by wires on the anodes. Light is observed by APDs behind the wires.

Energy from Ionization and Scintillation



Energy is independently measured from scintillation and ionization.

They are anticorrelated – better energy resolution from both together than either independently.



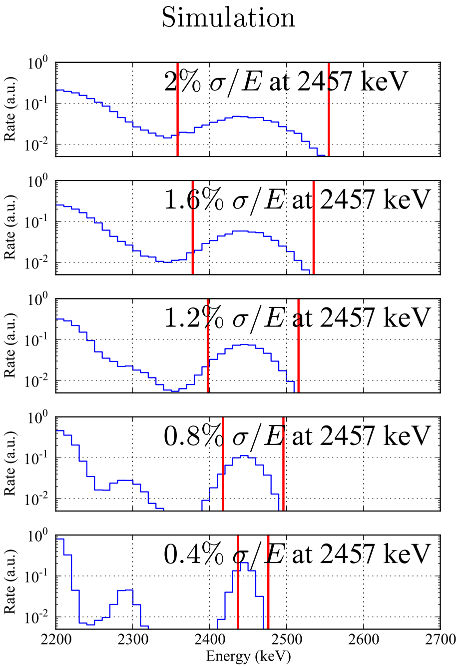
Primary Backgrounds: ^{232}Th, ^{238}U, and ^{137}Xe

Energy resolution is measured as σ/mean of a peak at the Q-value. Typically 1.5-2% for us.

Impact of background goes like integral of pdf between red 2σ lines.

^{232}Th gamma line at 2615 keV, so energy resolution has strong impact down to about 1.2%.

Beyond that, impact on ^{232}Th is much less.

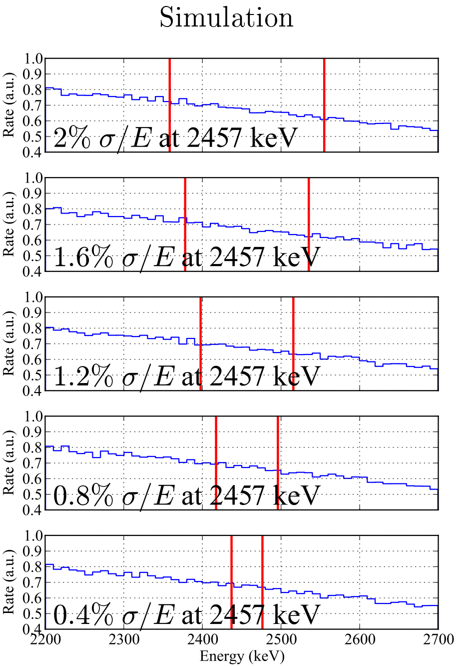


Primary Backgrounds: ^{232}Th, ^{238}U, and ^{137}Xe

^{238}U has a 2448-keV gamma line, indistinguishable from 2457-keV Q-value except with extremely good resolution.

So, even down to 0.4% energy resolution, most of the ^{238}U peak at 2448 keV is still within our energy window. Impact of resolution on ^{238}U backgrounds is small.

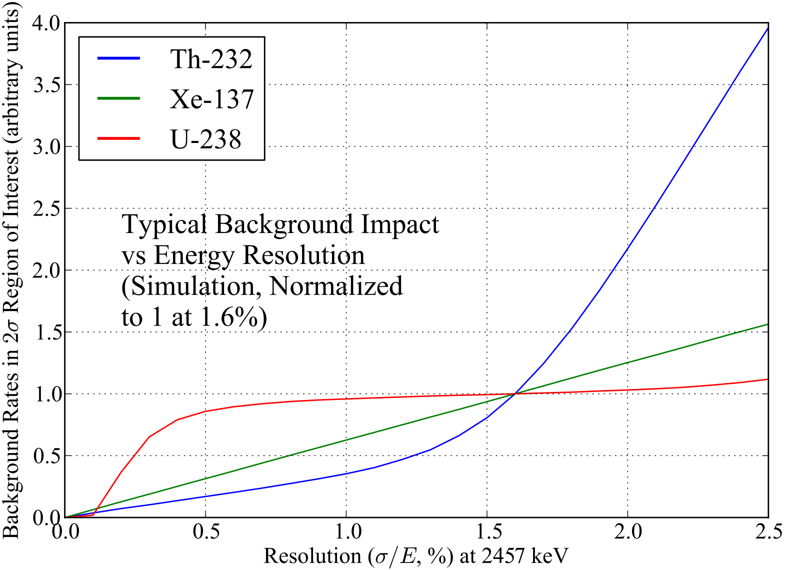
Primary Backgrounds: ^{232}Th, ^{238}U, and ^{137}Xe



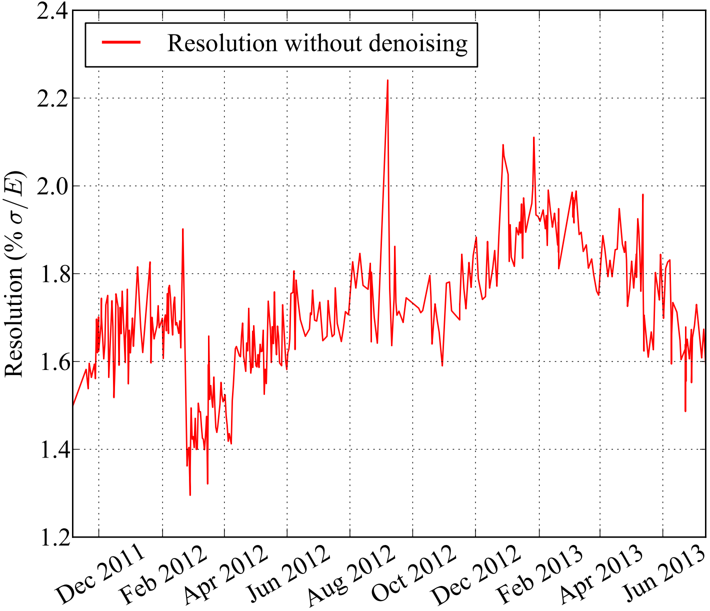
^{137}Xe is a cosmogenic background from ^{136}Xe which β-decays; endpoint of 4173 keV.

^{137}Xe spectrum is smooth around Q-value, so background impact goes linearly with resolution.

Backgrounds vs. Resolution



Time Variation of Resolution



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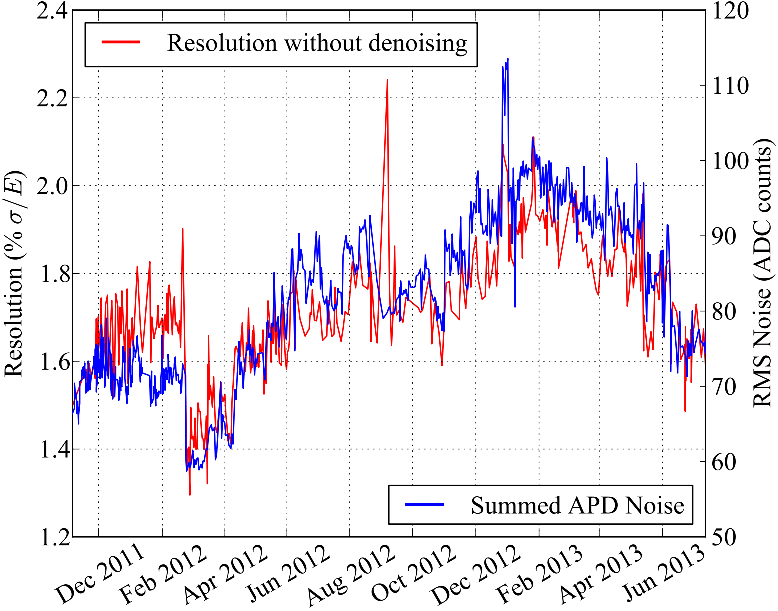
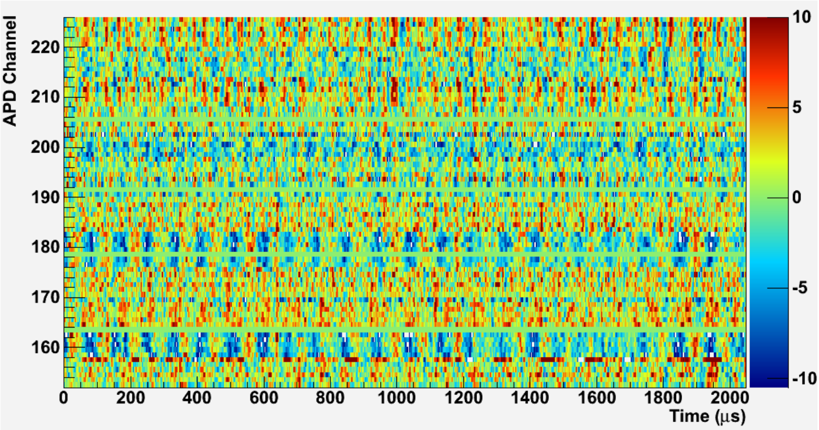


APD Noise is Correlated across Channels



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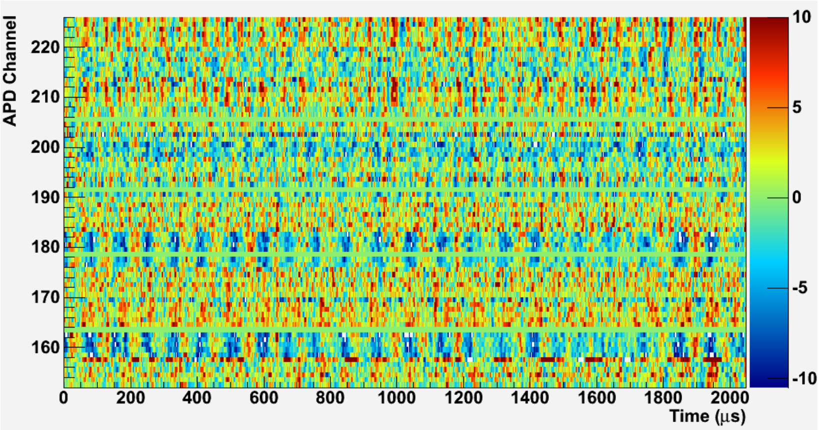

Correlated noise ⇒ offline denoising of some sort should work!

Types of Noise

Three types of noise in the scintillation measurements:

- Electronic noise.
- Photon fluctuations.
- Gain fluctuations.

A “denoising” algorithm should minimize all three.

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Most denoising algorithms transform an input waveform to an output waveform, with pulses amplified and additive noise reduced.

Here we're trying to reduce the impact of noise which is correlated with pulses. That means traditional denoising won't accomplish what we need.

So, "denoising" for us means producing a noise-tolerant estimate of the scintillation energy.

Types of Denoising

There are also three approaches to reducing the impact of noise:

- Frequency weighting.
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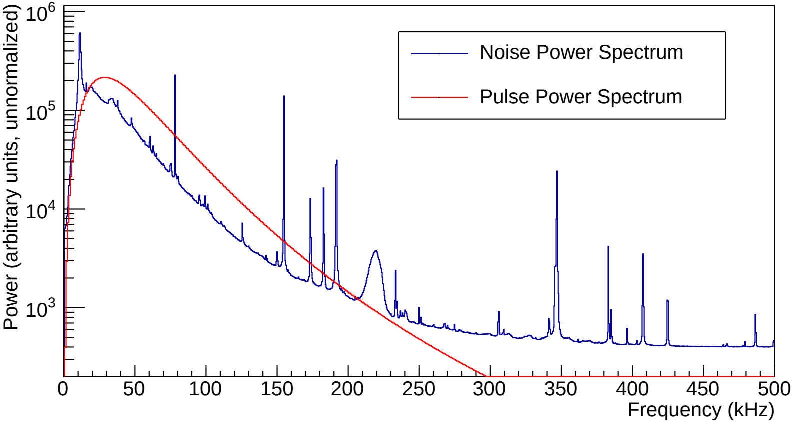
Noise and pulse power spectra.



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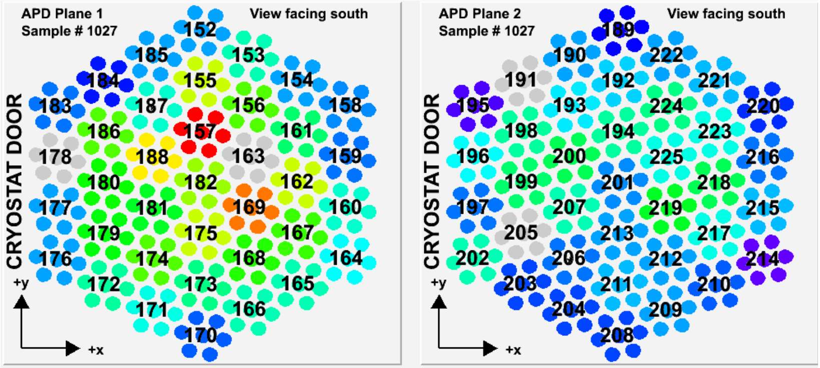
APD pulse magnitudes.



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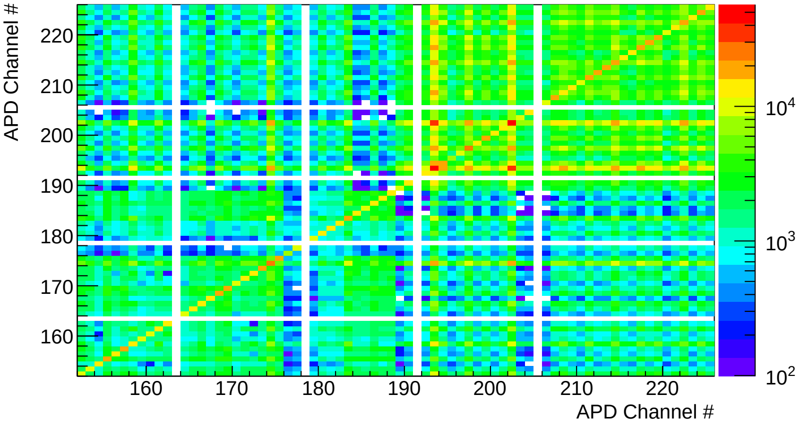
Noise correlations at 78.6 kHz.



Waveform Model

When there is one energy deposit in the xenon, the APD waveform $X_i[f]$ (in frequency space) on channel i can be modeled:

$$X_i[f] = M_i Y_i[f] + N_i[f], \text{ where:}$$

- $Y_i[f]$ is the unit-magnitude template pulse for channel i.
- M_i is the magnitude of the pulse observed on channel i, including the scintillation energy and Poisson fluctuations in photon statistics and gain.
- $N_i[f]$ is the additive (electronic) noise on channel i.

To complete our model, we need to understand the distributions of the random variables:

- How M_i depends on the unknown energy E.
- How magnitudes M_i and M_j are correlated.
- How electronic noise $N_i[f]$ and $N_j[f]$ are correlated.

Electronic Noise

To measure the electronic (additive) noise, we need waveforms with no pulse.
Fortunately, EXO-200 is a low-background detector: most of the time all it measures is noise.

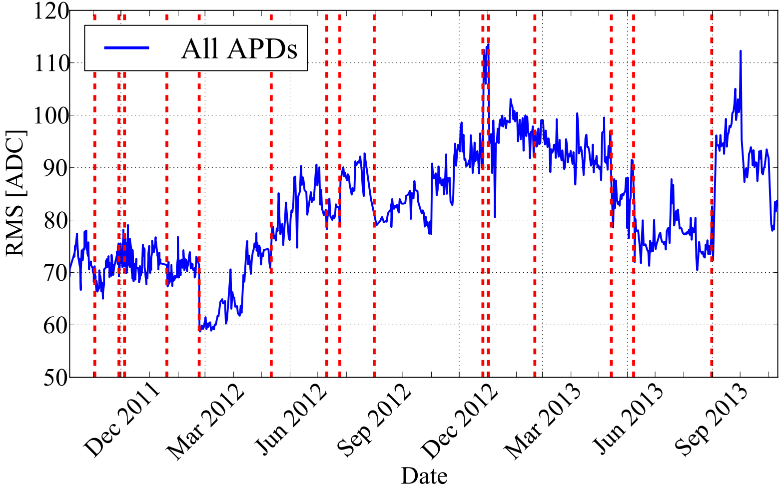
So, this is fairly easy: use our noise data to measure all of the pairwise noise correlations.
Explicitly, we measure:

$$\left\langle \mathrm{N}_i^R[f]\mathrm{N}_j^R[f] \right\rangle, \left\langle \mathrm{N}_i^R[f]\mathrm{N}_j^I[f] \right\rangle, \text{ and } \left\langle \mathrm{N}_i^I[f]\mathrm{N}_j^I[f] \right\rangle,$$

where $\mathrm{N}_i^R[f]$ and $\mathrm{N}_i^I[f]$ are the real and imaginary parts of $\mathrm{N}_i[f]$.

Electronic Noise

Main detail: the electronic noise changes over time (mostly in discontinuous steps).

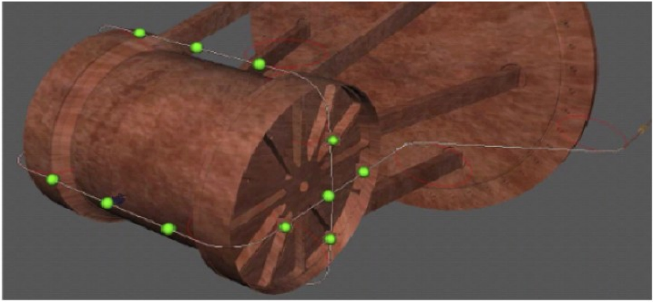


Lightmaps

The expected pulse height M_i on a channel i from a deposit at position $\vec{x}$ and time t is described by a lightmap $L_i(\vec{x}, t)$:

$$\langle M_i \rangle = L_i(\vec{x}, t)E.$$

To measure $L_i(\vec{x}, t)$, we need a known-energy deposit for all position and time bins. But we don't have that much data: calibration sources give us good statistics, but aren't near every position.



Lightmaps

The expected pulse height $\mathrm{M_i}$ on a channel i from a deposit at position $\vec{x}$ and time t is described by a lightmap $\mathrm{L_i}(\vec{x}, t)$:

$$\langle \mathrm{M_i} \rangle = \mathrm{L_i}(\vec{x}, t) E.$$

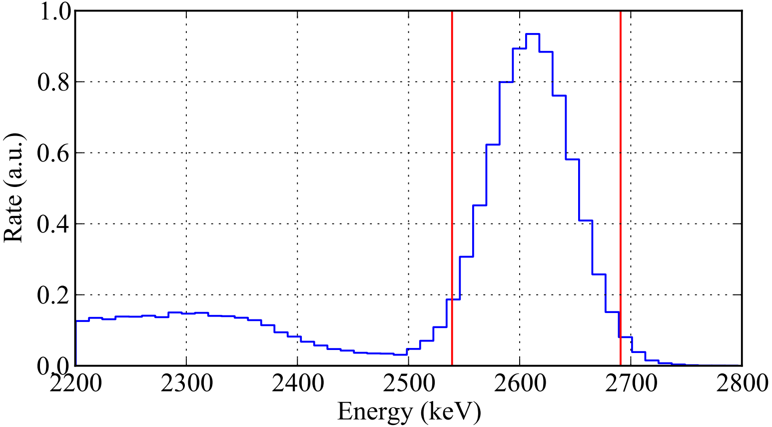
To reduce the amount of statistics needed, we make an approximation that $\mathrm{L_i}(\vec{x}, t)$ is separable into spatial and temporal components:

$$\mathrm{L_i}(\vec{x}, t) = \mathrm{R_i}(\vec{x}) \mathrm{S_i}(t).$$

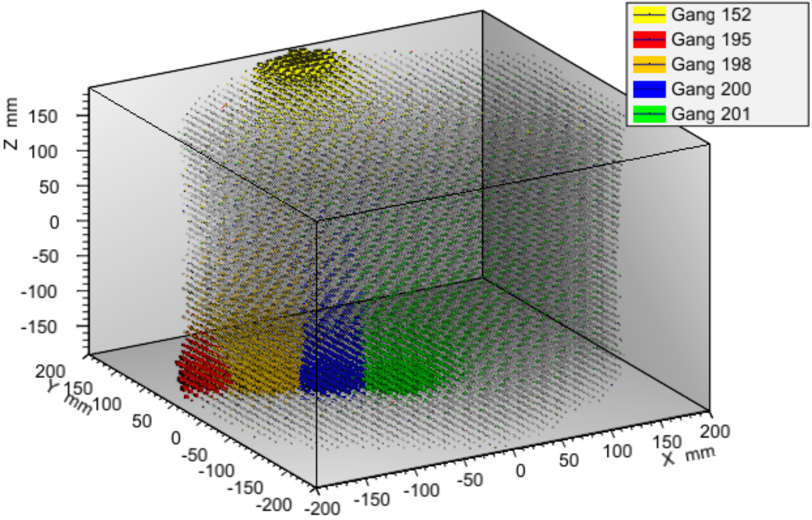
So, rather than needing one known-energy deposit per position per time, we need one per position during the whole history of EXO-200.
(We do this assuming that APD gain varies over time, but light collection does not, and gain does not depend on the position of the deposit.)

Lightmaps

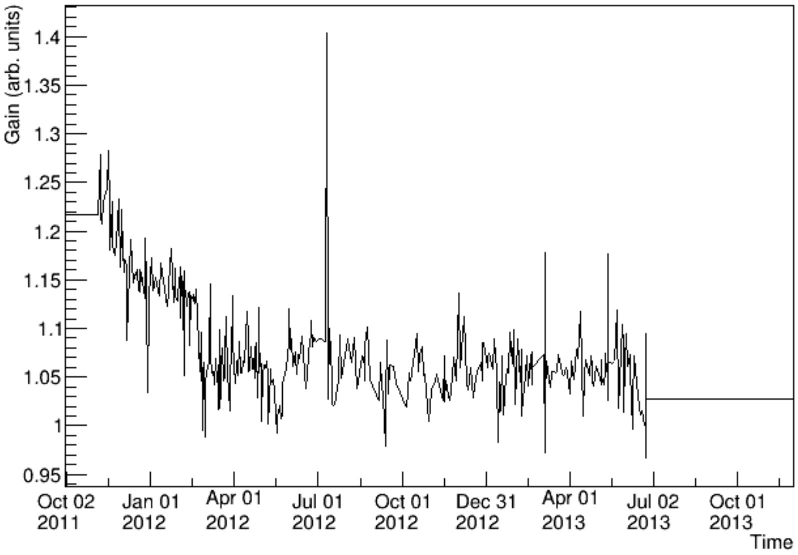
With this simpler form for the lightmap, we can measure $R_i(\vec{x})$ and $S_i(t)$ by combining all thorium source calibration events from the ^{208}Tl 2615-keV gamma line. It is a clean, well-isolated peak:



Lightmaps: $R_i(\vec{x})$



Lightmaps: $S_i(t)$ for $i = 152$



Light Collection Noise

The expected pulse height M_i on a channel i from a deposit at position $\vec{x}$ and time t is described by a lightmap $L_i(\vec{x}, t)$:

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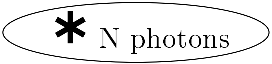
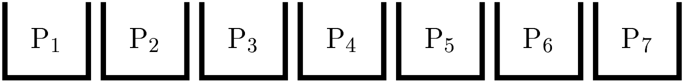
What about covariances of M_i?

Photon collection by the APDs has a multinomial distribution:



With photon collection $\langle P_i \rangle = N \cdot f_i$,
covariances are $\mathrm{cov}(P_i, P_j) = N \cdot \left(f_i \delta_{ij} - f_i f_j\right)$.



Photons randomly deposit on APDs (or are not observed).

Light Collection Noise

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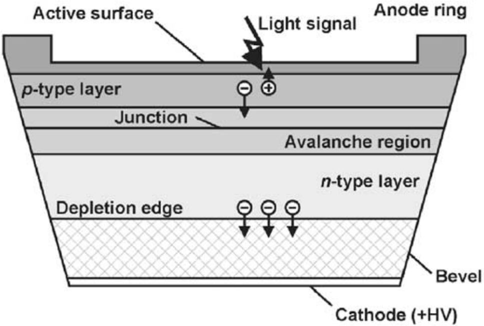
What about covariances of M_i?

We treat an APD as having two operational phases:



Moszyski et al., NIMA 2002.

- Active surface: photons are converted to photoelectrons (with a quantum efficiency).
- Avalanche region: photoelectrons are amplified by avalanche (APD gain).

Both processes produce Poissonian fluctuations.

Denoising: An Optimization Problem

So, the waveforms $X_i[f]$ are modeled by:

$$X_i[f] = M_i Y_i[f] + N_i[f],$$

and we have full descriptions of the random variables M_i and $N_i[f]$.
Linear real-valued energy estimators $\widehat{E}$ take the form:

$$\widehat{E} = \sum_{if} A_i[f] X_i^R[f] + B_i[f] X_i^I[f],$$

where $A_i[f]$ and $B_i[f]$ are parameters we can adjust.

Problem: minimize the mean square error, $\left\langle \left(E - \widehat{E}\right)^2 \right\rangle$.

Constraint: estimator is unbiased, $\left\langle E - \widehat{E} \right\rangle = 0$.

Denoising: An Optimization Problem

The solution is a big matrix equation

Denoising: An Optimization Problem

The solution is a big matrix equation which we write in a compact way:

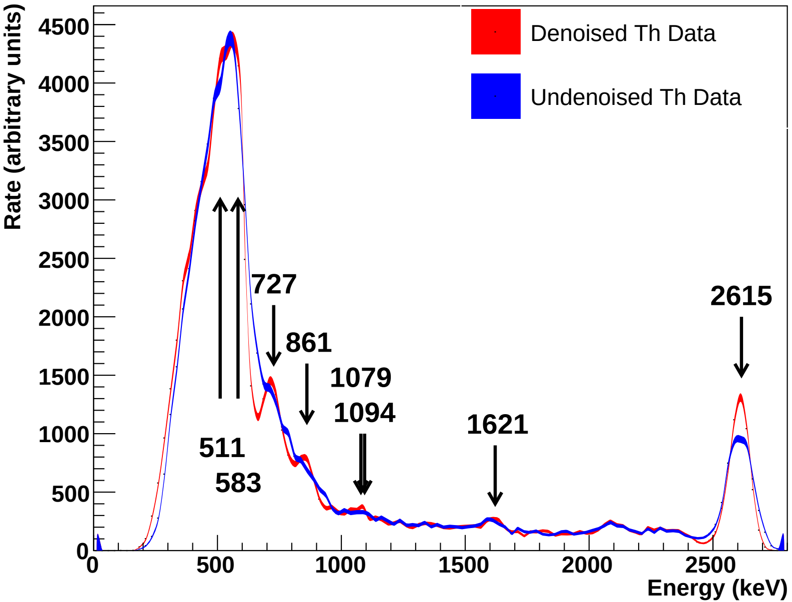
$$\begin{pmatrix} \mathrm{N}+\mathrm{P} & \mathrm{C}^{\top} \\ \mathrm{C} & 0 \end{pmatrix} \begin{pmatrix} \mathrm{A} \\ \lambda \end{pmatrix} = \begin{pmatrix} 0 \\ 1 \end{pmatrix}.$$

The LHS matrix is a $143{,}291 \times 143{,}291$ sparse matrix with roughly twenty million non-zero entries.

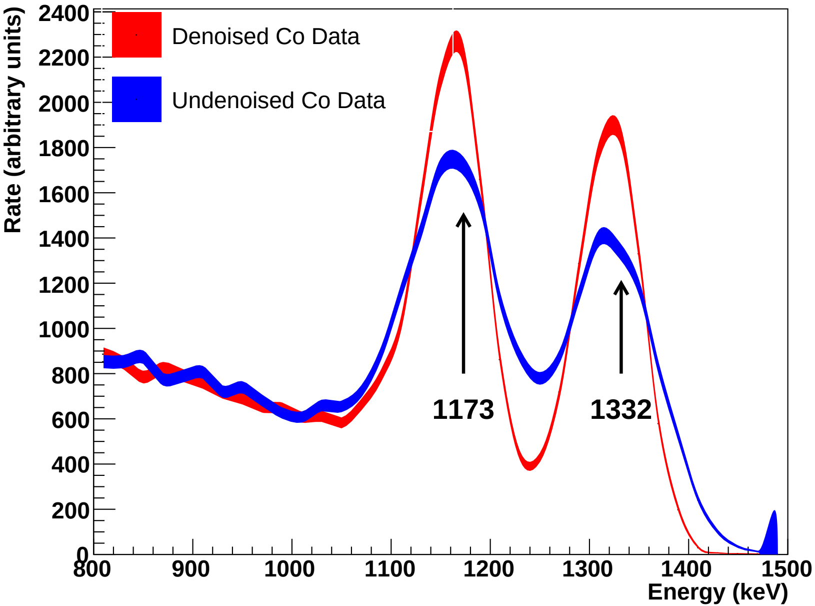
We need to solve this system once for each event, and there are hundreds of millions of events. Fortunately this is an embarassingly parallel problem, so we can just find a big computing system and do it.

NERSC, at LBNL, was used for this analysis; we required roughly 50,000 core-hours to do this processing.

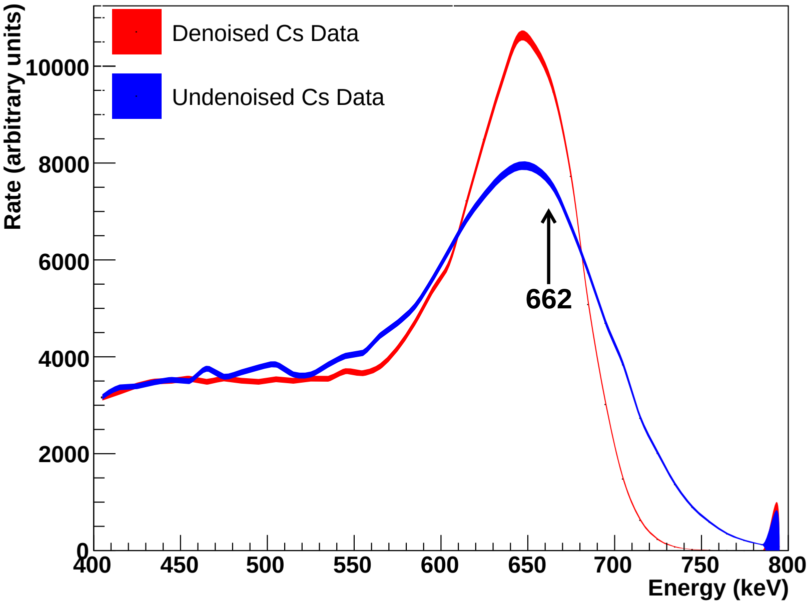
^{228}Th Source Spectrum (Before and After)



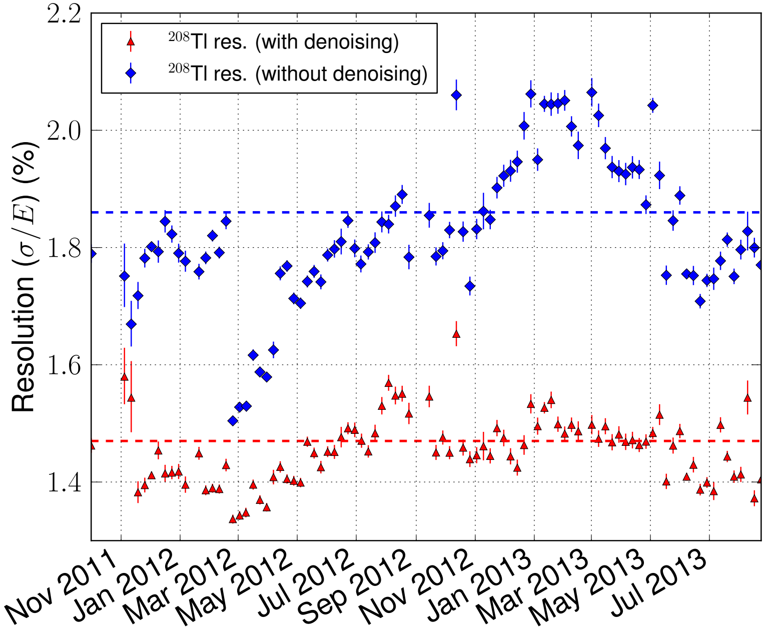
^{60}Co Source Spectrum (Before and After)



^{137}Cs Source Spectrum (Before and After)



Energy Resolution (Before and After)



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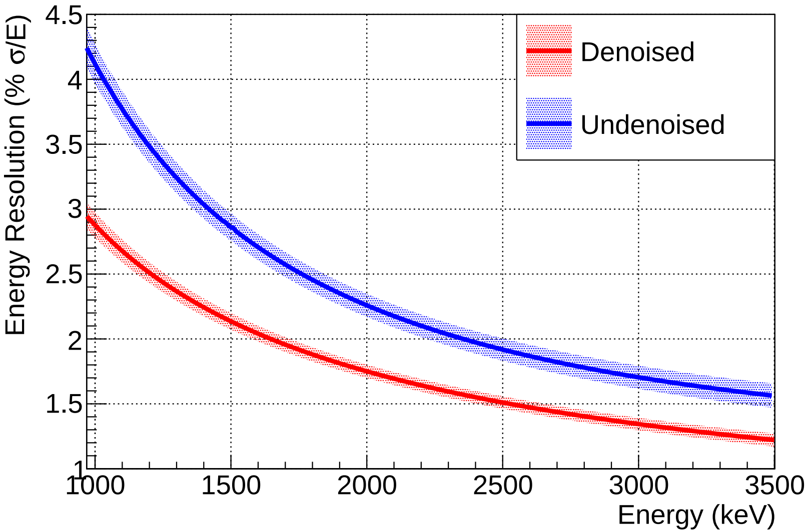
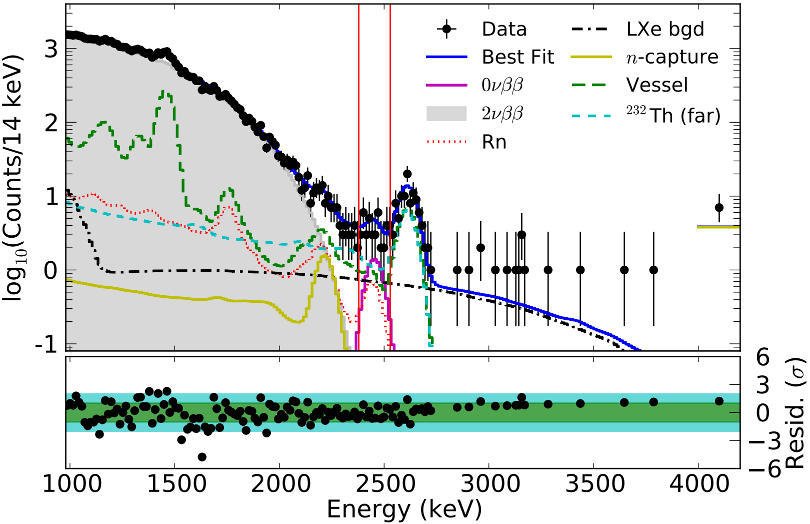
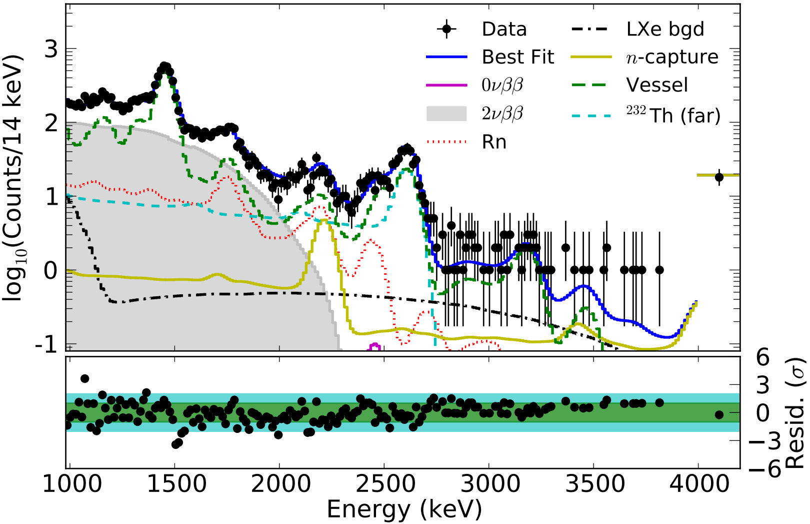


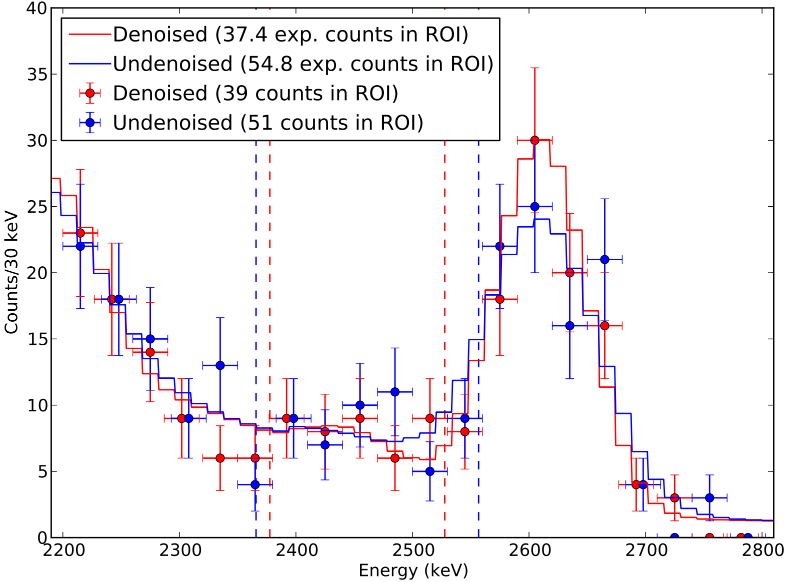
Best Fit (Denoised)



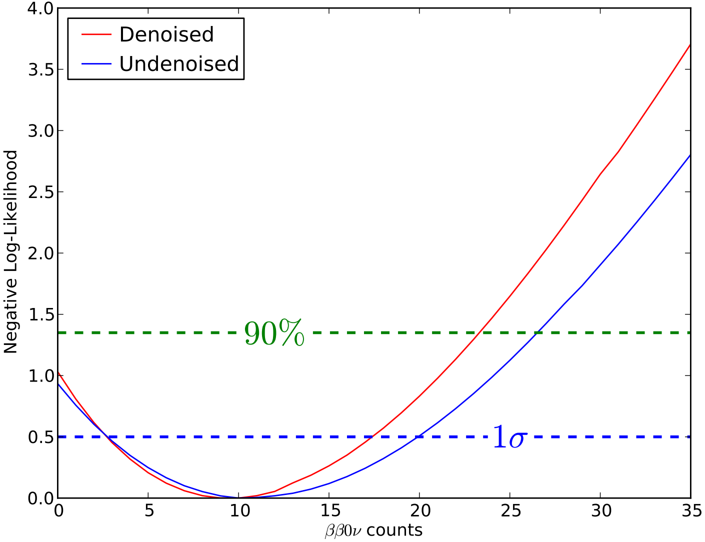
Best Fit (Denoised), multi-site backgrounds

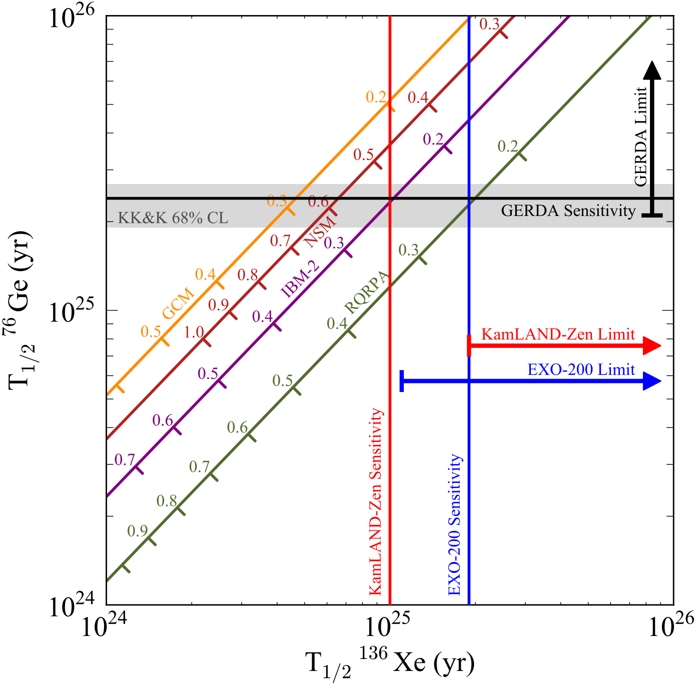


Fits around $\beta\beta 0\nu$ (Before and After)



Profile Likelihood (Before and After)





GERDA Limit
GERDA Sensitivity
KK&K 68% CL
GCM
NSM
IBM-2
RQRPA
KamLAND-Zen Limit
EXO-200 Limit
KamLAND-Zen Sensitivity
EXO-200 Sensitivity
$T_{1/2}$ ^{76}Ge (yr)
$T_{1/2}$ ^{136}Xe (yr)
10^{24}
10^{25}
10^{26}

The EXO-200 Collaboration

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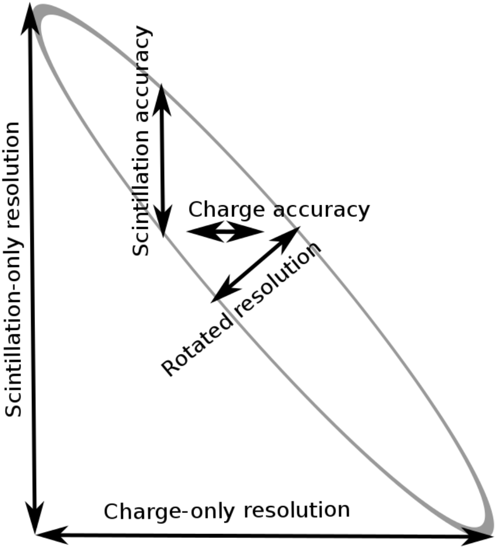
Thank You!

Questions?

Backup Slides

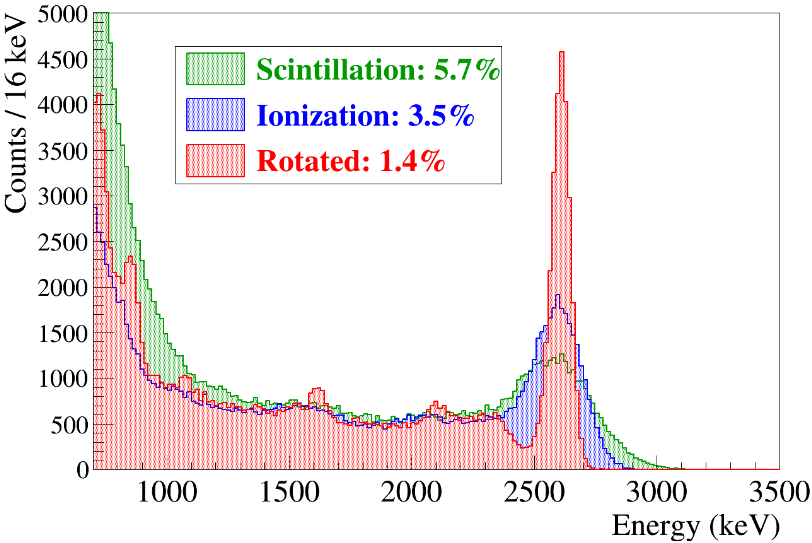
Anticorrelated Scintillation/Charge



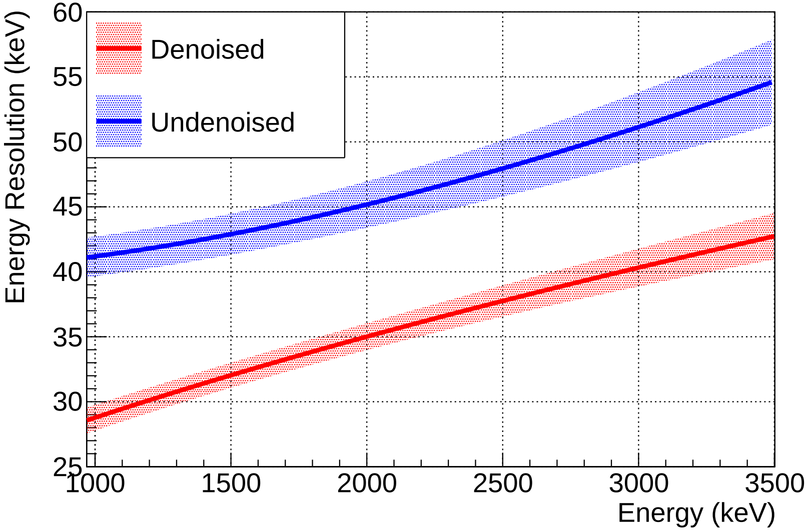
Why not the Anscombe transformation?

The Anscombe transformation (and generalizations) are designed to transform Poisson noise into Gaussian noise. However, for this purpose "noise" is defined to mean noise uncorrelated with pulse. We, on the other hand, are looking to understand noise which is correlated with the pulse, and then minimize its effect.
So, the Anscombe transformation doesn't really help us here.

Impact of Rotated Energy



Energy Resolution (Before and After)



Mean Sensitivity