

DESI: The PRObabilistic Value-Added Bright Galaxy Survey (PROVA-BGS) Mock Challenge

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ABSTRACT

In this project, we present the methodology for inferring physical properties of galaxies from joint SED fitting of the DESI optical photometry and spectroscopy. We construct realistic forward modeled DESI data from star formation and chemical enrichment histories of galaxies in the L-GALAXIES semi-analytic model and the ILLUSTRIS TNG hydrodynamic simulations. Then, using these mock observations, we demonstrate that the stellar mass and SFR posteriors from our SED fitting are consistent with constraints from other SED fitting methods in the literature and, more importantly, the input stellar masses and SFRs from the simulations. The SED fitting we present and validate in this project will be used to construct the PRObabilistic Value-Added Bright Galaxy Survey (PROVABGS) from DESI observations.

Keywords: keyword1 – keyword2 – keyword3

1. INTRODUCTION

What is DESI? Provide an overview of DESI specifics, numbers, and science goals, which will mostly be cosmology (BAO, RSD, etc). DESI will be great

TODO

In addition to the impact DESI will have on cosmology, DESI will also provide a transformative data set for galaxy science. **emphasize DESI samples useful for galaxy science: magnitude complete sample of BGS, LRGs, QSOs.** With this statistically powerful sample of, **Brief list of some exciting galaxy and quasar physics you can do with the DEIS BGS, QSOs, etc samples.**

TODO
 TODO
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Current status of DESI

- Comissioning completed
- SV data by the end of the year

As we prepare for the imminent torrent of data, in this paper we present the timely mock challenge of the DESI Galaxy Quasar Physics (hereafter GQP) working group. **what is the mock challenge?**

TODO
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Why do we need a mock challenge? We want to test and cement our methodology specifically for our GQP analysis before SV data comes out. As part of the survey preparation, we have all the tools

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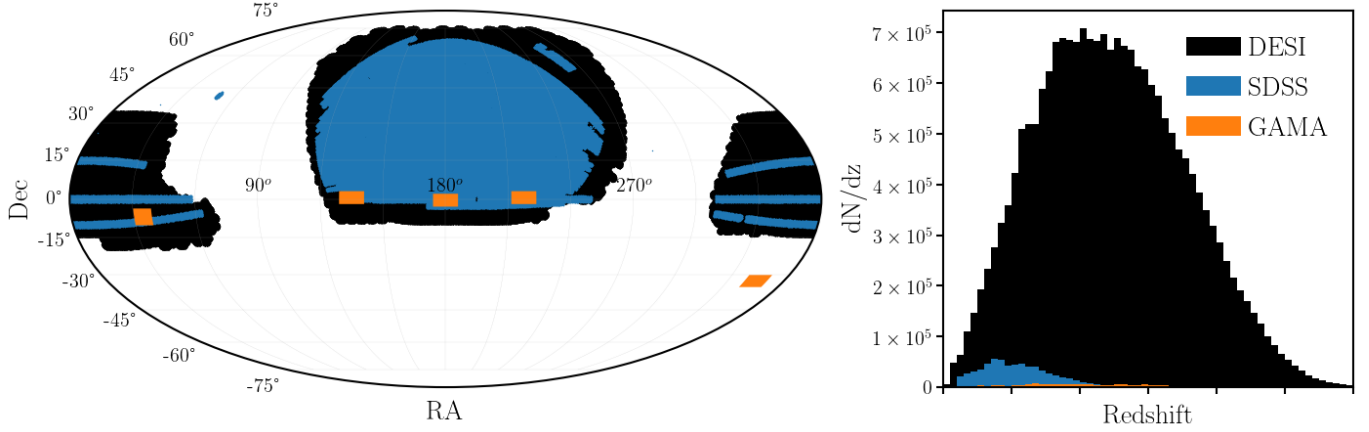


Figure 1. DESI will conduct the largest spectroscopic survey to date covering $\sim 14,000 \text{ deg}^2$. During dark time, DESI will measure >20 million spectra of luminous red galaxies, emission line galaxies, and quasars out to $z > 3$. During bright time, DESI will measure the spectra of ~ 10 million galaxies out to $z \sim 0.4$ with the Bright Galaxy Survey (BGS). *Left:* With its $\sim 14,000 \text{ deg}^2$ footprint (black), DESI will cover $\sim 2\times$ the SDSS footprint (blue; $\sim 7000 \text{ deg}^2$) and $\sim 45\times$ the GAMA footprint (orange; $\sim 300 \text{ deg}^2$). *Right:* Over this footprint, the BGS will provide spectra for a magnitude limited sample of ~ 10 million galaxies down to $r < 20$, two orders of magnitude deeper than the SDSS main galaxy sample and 0.2 mag deeper than GAMA.

to accurately forward model observations. **details on some of the specific tools and what we're able to simulate: realistic spectroscopy. realistic photometry. realistic spectro-photometry** All of this gives us a rare opportunity to test our methodology on bespoke simulations.

A mock challenge is also great for testing new methodology. BGS is a bright time survey and will push the boundaries of low SNR spectra. But if we can find a way to infer robust galaxy properties the statistical payout is awesome. **Something also about LRGs** We're also trying to robustly fit spectra and photometry simultaneously. This has been done before (**citations**) but not extensively tested on simulations.

broader impact of mock challenge highlight some gqp science cases with a focus on mock challenge papers

2. SIMULATIONS

DESI, with its robotically-actuated and fibre-fed spectrographs, will collect 5000 spectra simultaneously. The spectra cover the wavelength range 3600 to 9800 \AA , with a spectral resolution of $R = \lambda/\Delta\lambda$ between 2000 and 5500. During its five-year operation, starting in 2020, it will measure over 30 million spectra over $14,000 \text{ deg}^2$ of the sky (**cite**). In addition, these DESI targets also have optical and infrared imaging data from the DESI Legacy Imaging Surveys (hereafter Legacy Surveys ?). The Legacy Surveys are a combination of three public projects (Dark Energy Camera Legacy Survey, Beijing-Arizona Sky Survey, and Mayall z -band Legacy Survey) that jointly imaged the $\sim 14,000 \text{ deg}^2$ DESI footprint in three optical bands (g , r , and z). Furthermore, DR8 of the Legacy Survey also includes photometry in the WISE $W1$, $W2$, $W3$, and $W4$ infrared bands. **The infrared**

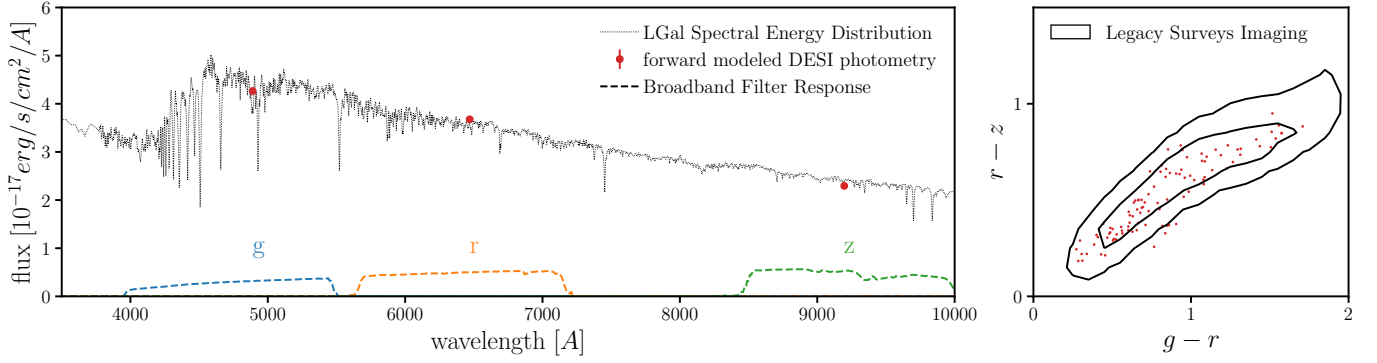


Figure 2. *Left:* We forward model DESI photometry (red) for our simulated galaxies (Section ??) by convolving their SEDs with the broadband filters (dashed). We construct the SEDs **CH: using the star formation and metallicity histories** (Section ??). We assign uncertainties to the forward modeled photometry by **CH: matching the colors to the Legacy Survey imaging for BGS** (Section ??). *Right:* The $g-r$ and $r-z$ color distribution of the forward modeled LGal photometry is consistent with the color distribution of BGS targets from Legacy Survey imaging (black contours).

photometry is from all imaging through year 4 of NEOWISE-Reactivation force-photometered in the unWISE maps at the locations of Legacy Surveys optical sources (cite). Below we describe how we simulate realistic DESI-like galaxy spectra and photometry from state-of-the-art simulations.

2.1. LGal

@rita

overview of Lgal. small volume semi-analytic and cosmo-hydro simulations

describe what galaxy properties (SFH, ZH, etc) are available

TODO

2.2. Spectral Energy Distributions

describe how the SED is generated using the SFH and ZHs

TODO

2.3. Forward Modeled DESI Photometry

From the noiseless source spectra generated for the galaxies in our simulations (Section ??), we generate noiseless photometry by convolving the spectra with the broadband filters of DESI imaging:

$$f_X = \int f(\lambda) R_X(\lambda) d\lambda \quad (1)$$

$f(\lambda)$, R_X is the transmission.

Fig. ??

Next, we assign realistic photometric uncertainties to the simulated photometric fluxes using the

Afterwards, we impose the target selection criteria of BGS: $r < 20$. This leaves us with **X** LGal galaxies.

TODO

2.4. Forward Modeled DESI Spectroscopy

fiber flux assigned when we assigned photometric uncertainty

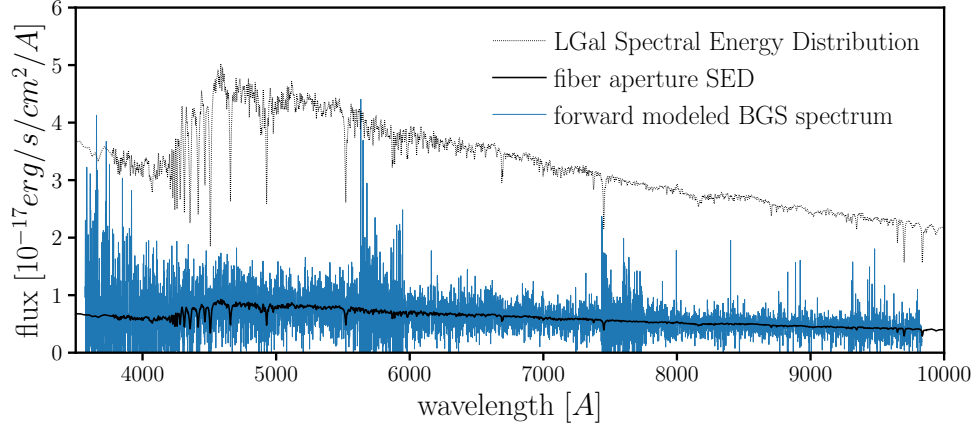


Figure 3. We construct forward model DESI spectra (blue) for our simulated galaxies by first applying a fiber aperture correction to the SED (black solid), then applying a DESI noise model. We apply a fiber aperture correction by scaling down the SED based on the r -band aperture flux we assign from the Legacy Surveys imaging. The noise model includes a DESI instrument model, which accounts for the DESI spectrograph response, and an atmosphere model, which accounts for the bright time observing conditions. We construct simulated DESI spectra for 8 different bright time observing conditions. For more details, we refer readers to Section ??.

- 'true' fiber aperture SED generated by scaling down the SED by the assigned r -band fiber flux.
- Meanwhile we add some noise to the fiber flux to have 'measured' fiber flux.
- Furthermore, we later jointly fit photometry and spectra we marginalize over f_{fiber} .
- run true fiber aperture SED through the DESI noise model to produce DESI-like spectra.
- describe the noise model in detail
- ¹

A paragraph on how we combine these to get the spectrophotometry with figure **figure** showing the spectra and photometry on top of each other

3. INFERRING GALAXY PROPERTIES FROM PHOTOMETRY AND SPECTRA

- overview of SED fitting
- refer to appendix about different set ups
- be explicit about what physical properties are inferred from our SED fitting: stellar masses, ages, metallicities, star-formation histories. Explain how we derive SFRs from the SFHs (i.e. our choice of 100Myr)

¹ <https://specsimpl.readthedocs.io>

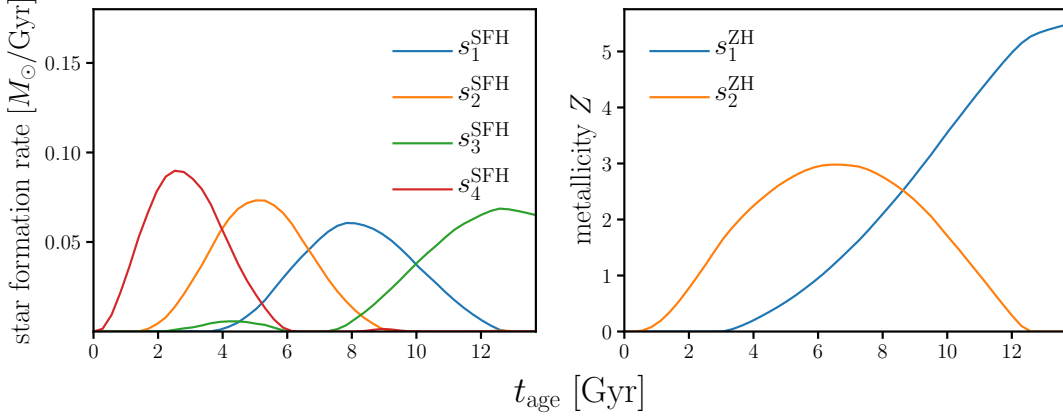


Figure 4. speculator bases

- brief description of other standard SED fitting methods we include for comparison: firefly, VESPA, CIGALE

description of our speculator SED model (?), which is based on FSPS. We use Chabrier IMF **CH:** do we need to justify htis?.

We don't use tau model for SFH because it biases the physical parameters inferred. citation. **TODO**

$$\text{SFH}(t, t_{\text{age}}) = \sum_{i=1}^4 \beta_i^{\text{SFH}} \frac{s_i^{\text{SFH}}(t)}{\int_0^{t_{\text{age}}} s_i^{\text{SFH}}(t) dt} \quad (2)$$

Similarly for ZH, we use the following ZH bases

$$\text{ZH}(t) = \sum_{i=1}^2 \gamma_i^{\text{ZH}} s_i^{\text{ZH}}(t) \quad (3)$$

$$\text{SFR}_{100\text{Myr}} = \frac{\int_{t_{\text{age}}-100\text{Myr}}^{t_{\text{age}}} \text{SFH}(t) dt}{100 \text{ Myr}} \quad (4)$$

see Figure ??.

Calzetti dust attenuation curve

Rather than evaluating the SPS directly using FSPS we use SPECULATOR, an emulator for FSPS. Brief explanation of the PCA neural net. motivate why we're using speculator by reiterating ?s speed profiling.

describe our MCMC sampling. James's work in convergence here. talk about how in principle speculator can easily be used with HMC because you can get derivatives with backpropagation. provide detailed profiling of SED fitting and projections for full 10million galaxy BGS sample.

In Figure XXXX we present the posterior as well as the best-fit photometry and spectroscopy using speculator + emcee.

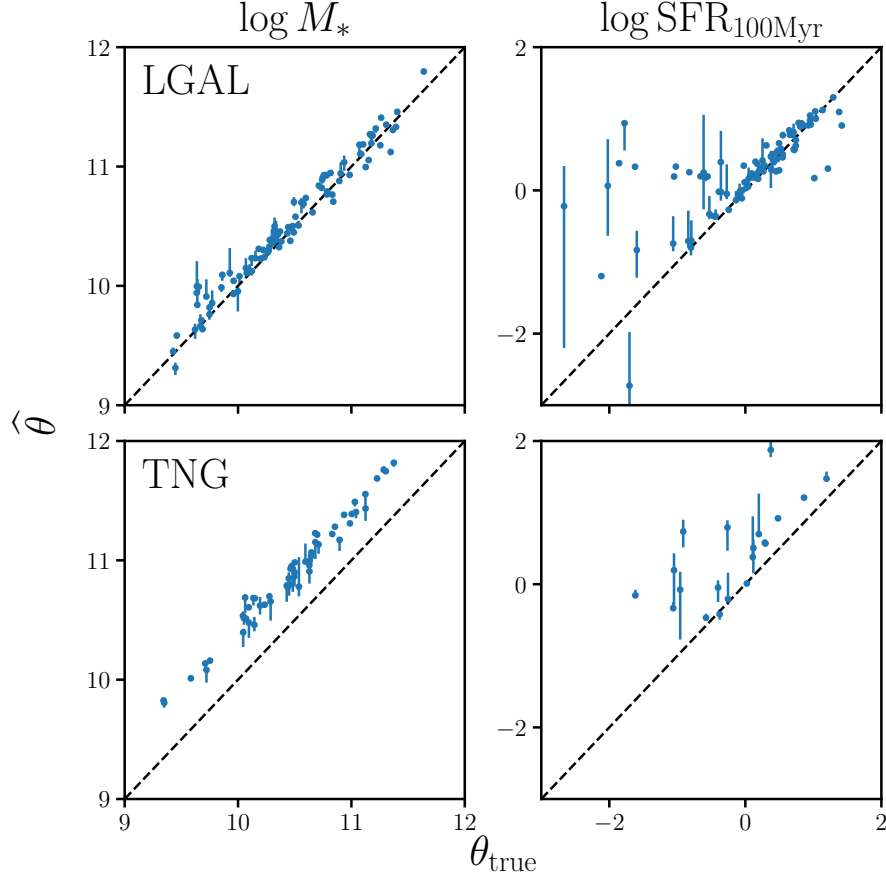


Figure 5. The properties inferred from ifspS spectrophotometry fit as a function of true properties.

4. RESULTS

Figure ??

In order to quantify the precision and accuracy of the inferred physical properties for our simulated galaxy population, we begin by assuming that the discrepancy between the inferred and true parameters for each galaxy $\Delta_{\theta,i}$

$$\theta_i^{\text{inf}} = \theta_i^{\text{true}} + \Delta_{\theta,i} \quad (5)$$

where $\Delta_{\theta,i}$ is sampled from a Gaussian distribution

$$\Delta_{\theta,i} \sim \mathcal{N}(\mu_{\Delta_\theta}, \sigma_{\Delta_\theta}). \quad (6)$$

This Gaussian distribution is described by population hyperparameters μ_{Δ_θ} and σ_{Δ_θ} , the mean and standard deviation, which quantify the accuracy and precision of the inferred physical properties for the population.

Given the photometry and spectrum of our galaxies, $\{\mathbf{D}_i\}$, we can get the posteriors for these population hyperparameters $\eta_\Delta = \{\mu_{\Delta\theta}, \sigma_{\Delta\theta}\}$ using a hierarchical Bayesian framework (?):

$$p(\eta_\Delta | \{\mathbf{D}_i\}) = \frac{p(\eta_\Delta) p(\{\mathbf{D}_i\} | \theta_\Delta)}{p(\{\mathbf{D}_i\})} \quad (7)$$

$$= \frac{p(\eta_\Delta)}{p(\{\mathbf{D}_i\})} \int p(\{\mathbf{D}_i\} | \{\theta_i\}) p(\{\theta_i\} | \eta_\Delta) d\{\theta_i\}. \quad (8)$$

Naively the posteriors for each of the galaxies are not correlated, so we can factorize the expression above

$$p(\eta_\Delta | \{\mathbf{D}_i\}) = \frac{p(\eta_\Delta)}{p(\{\mathbf{D}_i\})} \prod_{i=1}^N \int p(\mathbf{D}_i | \theta_i) p(\theta_i | \eta_\Delta) d\theta_i \quad (9)$$

$$= \frac{p(\eta_\Delta)}{p(\{\mathbf{D}_i\})} \prod_{i=1}^N \int \frac{p(\theta_i | \mathbf{D}_i) p(\mathbf{D}_i)}{p(\theta_i)} p(\theta_i | \eta_\Delta) d\theta_i \quad (10)$$

$$= p(\eta_\Delta) \prod_{i=1}^N \int \frac{p(\theta_i | \mathbf{D}_i) p(\theta_i | \eta_\Delta)}{p(\theta_i)} d\theta_i. \quad (11)$$

$p(\theta_i | \mathbf{D}_i)$ is the posterior for galaxy i . Hence, the integral can be which means the integral can be estimated using the MCMC sample from the posterior

$$p(\eta_\Delta | \{\mathbf{D}_i\}) = p(\eta_\Delta) \prod_{i=1}^N \frac{1}{S_i} \sum_{j=1}^{S_i} \frac{p(\theta_{i,j} | \eta_\Delta)}{p(\theta_{i,j})}. \quad (12)$$

S_i is the number of MCMC samples and $\theta_{i,j}$ is the j^{th} sample of galaxy i . We present the maximum a posteriori (MAP) estimates of η_Δ for $\log M_*$ and $\log \text{SFR}$ in Figure ??.

η_Δ as a function of SNR/mag/colour.

discussion:

- sfh basis improves SFH accuracy?
- advantage of mcmc over MAP
- non-Gaussian posteriors
- science applicatoinis
- comparison to other methods
- different observing condition doesn't impact our results?
- we use super simple dust mmodel. **CH:** more sophistiicated dust models doesn't matter?
- the simulated spectra and the fitting are generated using the same SED. **CH:** We try with different libraries.
- **CH:** we try fitting with a different IMF

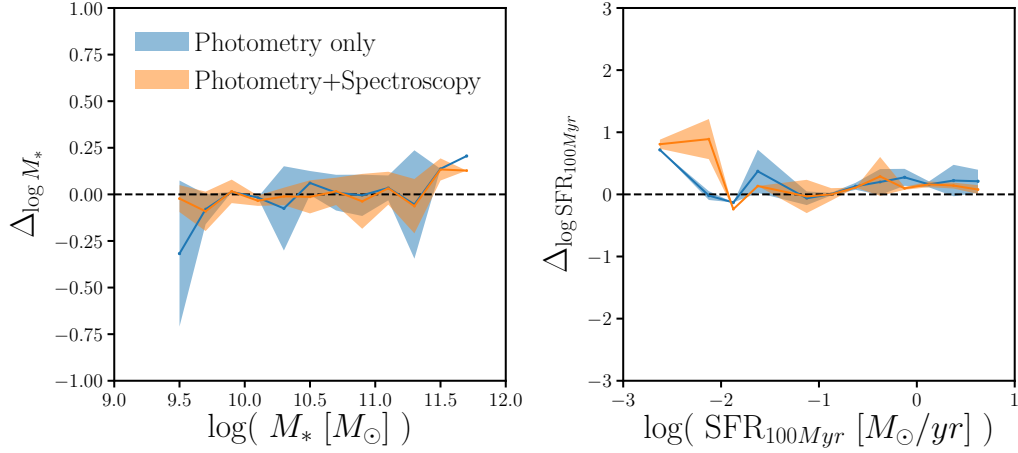


Figure 6. The discrepancies between the inferred and input/“true” M_* s (left) and SFRs (right) for our LGAL galaxies. In blue, we infer M_* s and SFRs using only photometry; in orange, we infer M_* s and SFRs by jointly fitting both photometry and spectroscopy. *Jointly fitting spectroscopy and photometry improves constraints on galaxy properties.*

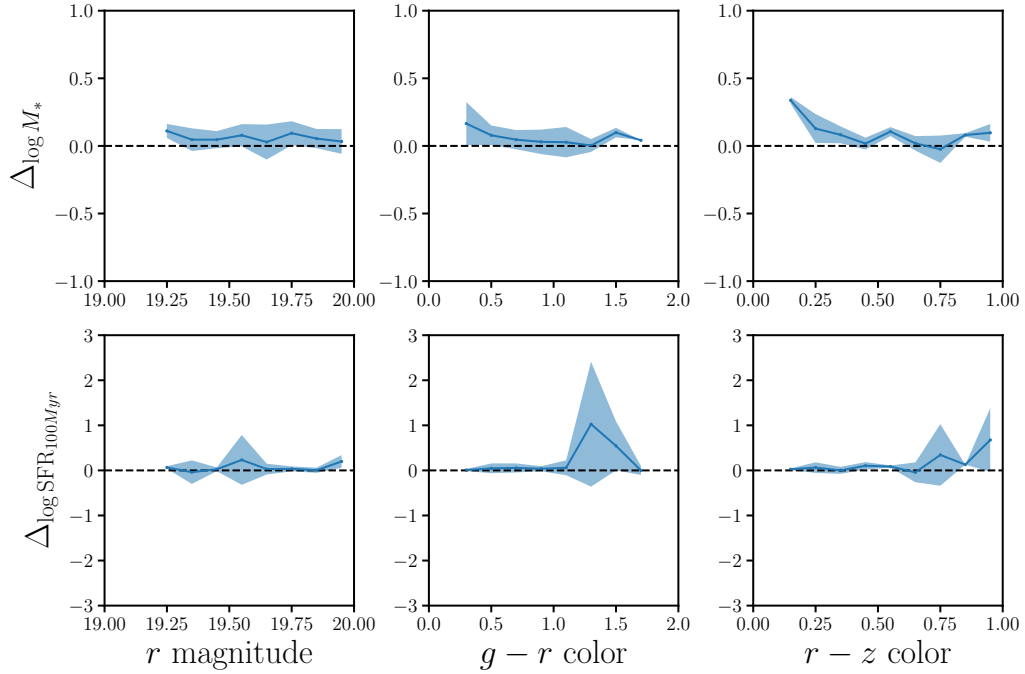


Figure 7. do we want to test bias against other properties? (e.g. obs condition?)

5. SUMMARY

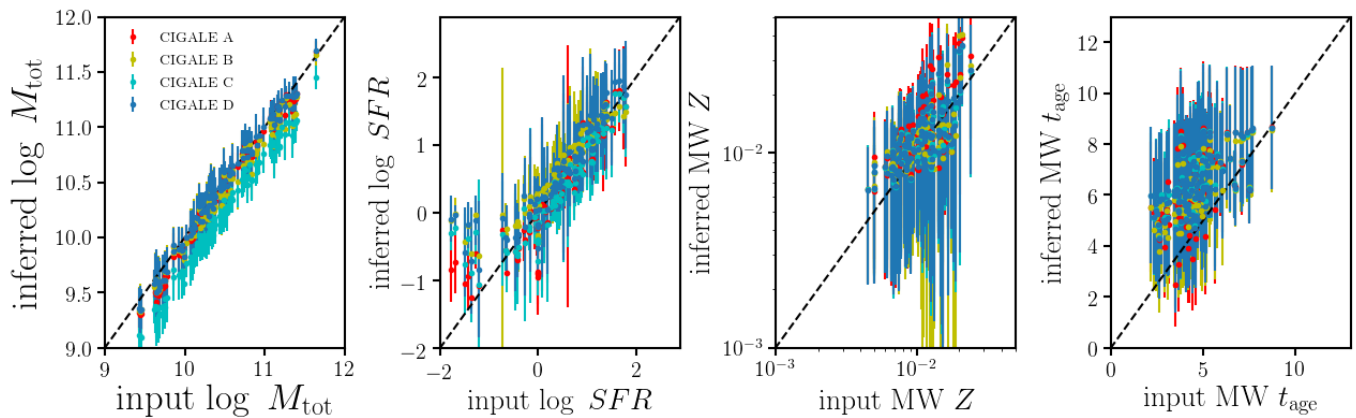
DESI is great.

Also advertise science papers with specific focus on mock challenge science papers.

ACKNOWLEDGEMENTS

Table 1. Table of $\theta_{\text{inf}} - \theta_{\text{true}}$ and uncertainties for different CIGALE setups

set	CIGALE A	CIGALE B	CIGALE C	CIGALE D
ΔM_{tot}	0.07	0.09	0.27	0.05
M_{err}	0.13	0.20	0.16	0.16
ΔAge	1.78	1.59	1.97	1.92
Age_{err}	2.36	2.63	2.47	2.47
ΔZ	0.0037	0.0026	0.0027	0.0027
Z_{err}	0.0091	0.0085	0.0089	0.0089

**Figure 8.** The properties inferred from CIGALE photometry fit as a function of true properties. Configuration CIGALE A, B, C, and D on one plot.

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APPENDIX

A. COMPARISON TO OTHER METHODS

REFERENCES

- Alsing J., et al., 2019, arXiv:1911.11778 [astro-ph] Hogg D. W., Myers A. D., Bovy J., 2010, [The Astrophysical Journal](#), 725, 2166
- Dey A., et al., 2019, [AJ](#), 157, 168