# k-Nearest Neighbor Estimator for Bayesian Evidence

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## ABSTRACT

abstract

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## 1. ESTIMATOR

Consider the commonly used Kullback-Leibler (KL) divergence between two distributions p(x) and g(x):

$$D_{\mathrm{KL}}(p || q) = \int p(x) \log \frac{p(x)}{q(x)} \, \mathrm{d}x. \tag{1}$$

The KL divergence between the posterior  $p=p(\theta\,|\,D,m)$  and prior  $\pi=p(\theta\,|\,m)$  is then

$$D_{\mathrm{KL}}(p \mid\mid \pi) = \int p(\theta \mid D, m) \log \frac{p(\theta \mid D, m)}{p(\theta \mid m)} d\theta.$$
 (2)

If we use Bayes' Theorem to substitute for  $p(\theta \mid D, m)$  in the numerator, we can rewrite this as

$$D_{\mathrm{KL}}(p \mid\mid \pi) = \int p(\theta \mid D, m) \log \frac{p(D \mid \theta, m)}{p(D \mid m)} d\theta$$
(3)

$$= -\log p(D \mid m) + \int p(\theta \mid D, m) \log p(D \mid \theta, m) d\theta.$$
 (4)

The first term of the right hand side is the log evidence and the second term is the expectation value of the log likelihood under the posterior,  $\langle \log \mathcal{L} \rangle$ .

In other words, if we can estimate the KL divergence between the posterior and prior we can also estimate the evidence:

$$p(D \mid m) = \langle \log \mathcal{L} \rangle - D_{KL} \left( p(D \mid \theta, m) \mid \mid \pi \right). \tag{5}$$

In standard Bayesian analyses using MCMC,  $\langle \log \mathcal{L} \rangle$  can be easily derived using samples from the posteriors and Monte Carlo integration.

Meanwhile, non-parametric divergence estimators also provide a way to exploit samples from the posterior to estimate the divergence. These estimators, which have been applied to Support

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Distribution Machines and used in the machine learning and astronomical literature (e.q. Póczos et al. 2011; Poczos et al. 2012a,b; Xu et al. 2013; Ntampaka et al. 2015, 2016; Ravanbakhsh et al. 2017; ?), allow us to estimate the divergence, D(p||q), using samples  $X_{1:n} = \{X_1, ... X_n\}$  and  $Y_{1:m} = \{Y_1, ... Y_m\}$ drawn from p and q respectively. In our case, p would be the posterior and  $X_{1:n}$  would be n samples drawn from the posterior (i.e. our MCMC chain) and  $Y_{1:m}$  would be m samples drawn from the prior. For the KL divergence, we use the k-Nearest Neighbor estimator presented in Wang et al. (2009):

 $D_{\mathrm{KL}}(p || q) \approx \widehat{D_{\mathrm{KL}}}(p || q) = \frac{d}{n} \sum_{i=1}^{n} \left[ \log \frac{\nu_{\ell_{i}}(i)}{\rho_{k_{i}}(i)} \right] + \frac{1}{n} \sum_{i=1}^{n} \left[ \psi(\ell_{i}) - \psi(k_{i}) \right] + \log \frac{m}{n-1}.$ 

$$D_{\mathrm{KL}}(p || q) \approx D_{\mathrm{KL}}(p || q) = \frac{\pi}{n} \sum_{i=1}^{n} \left[ \log \frac{\ell_i(\ell)}{\rho_{k_i}(i)} \right] + \frac{\pi}{n} \sum_{i=1}^{n} \left[ \psi(\ell_i) - \psi(k_i) \right] + \log \frac{\pi}{n-1}. \tag{6}$$
The largest the Euclidean distance of the  $k^{\mathrm{th}}$  nearest neighbor of  $X_i$  from sample  $X_i$  and  $\psi_i(i)$ 

 $\rho_k(i)$  denotes the Euclidean distance of the  $k^{\text{th}}$  nearest neighbor of  $X_i$  from sample  $X_{1:n}$  and  $\nu_\ell(i)$ denotes the Euclidean distance of the  $\ell^{\text{th}}$  nearest neighbor of  $Y_i$  in the sample  $Y_{1:m}$ .  $\psi$  is the Digamma function:

$$\psi(k) = \Gamma'(k)/\Gamma(k) \tag{7}$$

 $\ell_i$  (and  $k_i$ ) is the number of samples  $X_{1:n}$  or  $Y_{1:m}$  contained in the ball  $B(X_i, \epsilon(i))$  where  $\epsilon(i) = 1$  $\max(\rho(i), \nu(i))$ . The second term reduces the estimation bias that comes from nonuniformity of the distribution near each sample point. fill in the rest of the details for the estimator (e.g. asymptotic convergence of the estimators)

# TODO

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# **APPENDIX**

### REFERENCES

Ntampaka M., Trac H., Sutherland D. J., Battaglia N., Poczos B., Schneider J., 2015, The Astrophysical Journal, 803, 50

Ntampaka M., Trac H., Sutherland D. J., Fromenteau S., Poczos B., Schneider J., 2016, The Astrophysical Journal, 831, 135

Póczos B., Szabó Z., Schneider J., 2011, in 2011 19th European Signal Processing Conference. pp 1718–1722

Poczos B., Xiong L., Schneider J., 2012a, arXiv:1202.3758 [cs, stat]

Poczos B., Xiong L., Sutherland D., Schneider J., 2012b, Machine Learning Department

Ravanbakhsh S., Lanusse F., Mandelbaum R., Schneider J., Poczos B., 2017, in Thirty-First AAAI Conference on Artificial Intelligence.

Wang Q., Sanjeev K., Sergio V., 2009, IEEE TRANSACTIONS ON INFORMATION THEORY, 55, 2392

Xu X., Ho S., Trac H., Schneider J., Poczos B., Ntampaka M., 2013, The Astrophysical Journal, 772, 147