

PRIMUS: QUIESCENT FRACTION AS A FUNCTION OF ENVIRONMENT AND REDSHIFT

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ABSTRACT

We present the evolution of the quiescent fraction (f_Q) for galaxies in different density environments from $z = 0.0$ to 0.8 using spectroscopic redshift and multi-wavelength imaging data from PRISM Multi-object Survey (PRIMUS) and the Sloan Digital Sky Survey (SDSS). We construct a stellar mass limited galaxy population of ~ 40000 galaxies from PRIMUS within the redshift range $0.2 - 0.8$ and ~ 100000 galaxies from SDSS within redshift range $0.0375 - 0.145$. Using an evolving cut based on specific star formation, we classify the galaxies as quiescent or star-forming and measure its environment using a fixed cylindrical aperture method on a volume-limited *Environment Defining Population* we construct from PRIMUS and SDSS. Based on its environment we classify our target population into dense or sparse environments.

With the target galaxy population divided into quiescent and star-forming galaxies in dense and sparse environments, we calculate the stellar mass functions (SMFs) for each of these subsamples. With these SMFs, we compute the f_Q s to find that at $M_* \sim 10.5$, f_Q increases by ~ 0.15 from $z = 0.8$ to 0.0 for both high and low environments. In addition, throughout the redshift range the difference between the f_Q at high and low environment remains constant at < 0.1 . These results suggest that while f_Q increases for galaxies in both dense and sparse environments over redshift, the evolution of f_Q is independent of environment.

1. INTRODUCTION

Galaxies, in their detailed properties, carry the imprints of their surroundings, with a strong dependence of the quiescent fraction of galaxies on their local environment (e.g. Hubble 1936; Oemler 1974; Dressler 1980; Hermit et al. 1996; Guzzo et al. 1997; for a recent review see Blanton & Moustakas 2009). The strength of this dependence is itself a strongly decreasing function of galaxy stellar mass; at the extreme, the lowest masses ($< 10^9 M_\odot$) galaxies are quenched only in dense regions, and never in isolation (Geha et al. 2012). These effects vary with redshift at least in the densest clusters, as observed in the changing fraction of late-type spirals relative to the field found in studies of the morphology-density relation (Dressler 1984; Desai et al. 2007). Clearly understanding the properties of galaxies in the present-day universe requires a careful investigation of the role of environment, and how that role changes over time.

Nevertheless, the evolution of the role of environment is a relatively subtle effect and difficult to study. Although history of galaxies prior to $z \sim 1$ appears to have been one of rapid assembly, since that time the galaxy population has continued to evolve, but less dramatically. Although there are detectable changes in the population, the major classes of galaxies existed at $z \sim 1$, in roughly the same relative numbers as today (Bundy et al. 2006; Borch et al. 2006; Taylor et al. 2009; Moustakas et al. 2013a). Furthermore, at those redshifts we can also detect the dependence of galaxy properties on environment, with lower star-formation rate early-type galaxies populating the denser regions (Cooper et al. 2008; Patel et al. 2009; Kovač et al. 2010).

The most dramatic change in galaxy properties during the past eight billion years has been a remarkable decline in the star-formation rate of galaxies in the Uni-

verse (Hopkins & Beacom 2006). This decline appears dominated by decreases in the rates of star-formation of individual galaxies (Noeske et al. (2007)). There is evidence that a large fraction of the decline associated with strongly infrared-emitting starbursts (Bell et al. 2005; Magnelli et al. 2009). The decline does not appear to be due to the quenching of a large fractions of the star-forming population, as reflected in observations of the stellar mass function of quiescent and star-forming galaxies (Blanton et al. (2006), Bundy et al. 2006; Borch et al. 2006; Moustakas et al. 2013a). These findings leave little room for the participation of environmentally-driven quenching in the global census of star-formation. As Cooper et al. (2008) and others have pointed out, because the environmental dependence of total star-formation rates at fixed redshift is relatively small, environmentally effects are unlikely to cause the overall star-formation rate decline.

Thus, the impact of environment on galaxy formation has to be interpreted on top of the background of this overall decline affecting galaxies in all environments. The most straightforward investigation of would directly determine the star-forming properties of galaxies as a function of environment, stellar mass and redshift in a single, consistently analyzed data set. This analysis can reveal how galaxies are quenched in the universe over time, quantitatively establish the contribution of environmental effects to the overall trends, and reveal whether those trends happen equally in all environments. However, such an analysis has not been done previously due to the lack of sufficiently large samples. In this paper, we apply this approach using the Prism Multi-object Survey (PRIMUS; Coil et al. (2011), Cool et al. (2013)), the largest available redshift survey covering the epochs between $0 < z < 1$.

2. SAMPLE SELECTION

In this paper we are interested in measuring the evolution of the quiescent fraction over a wide range of redshifts and in different galaxy environments. In order to construct a sample with redshift depth and robust enough redshift values to measure galaxy environment, we use galaxies at intermediate redshifts from PRIMUS. Additionally we supplement our sample with galaxies at low redshift ($z \sim 0.1$) from SDSS.

We begin with a brief summary of the PRIMUS data in Section 2.1 followed by a summary of the SDSS data in Section 2.2. Then in Section 2.3 we use this data to define the stellar mass complete target galaxy population. Afterwards, in Section 2.4, we classify these target galaxies as quiescent or active star-forming galaxies. For each galaxy we obtain the environment using a volume-limited *Environment Defining Population* in Section 2.5. Finally in Section 2.6, we correct the galaxy sample and its environment measurements for the edge effects of the surveys.

2.1. PRIMUS

For galaxies at intermediate redshifts we use multi-wavelength imaging and spectroscopic redshifts data of PRIMUS, which is a faint galaxy survey with precise spectroscopic redshifts ($\sigma_z/(1+z) \approx 0.5\%$) for $\sim 120,000$ galaxies within redshifts $z \approx 0 - 1.2$. The survey was conducted using a IMACS spectrograph on a Magellan I Baade 6.5 m telescope with a slitmask and low dispersion prism. For further details on the PRIMUS observation methods, including survey design, targeting, and data summary, we refer readers to the survey papers: Coil et al. (2011) and Cool et al. (2013).

As done in Moustakas et al. (2013b), we only use fields targeted by PRIMUS with *GALEX* and *Spitzer*/IRAC imaging. This restricts us to five fields. Four of these fields are a part of the *Spitzer* Wide-area Infrared Extragalactic Survey (SWIRE¹): the European Large Area ISO Survey - South 1 field (ELAIS-S1²), the Chandra Deep Field South SWIRE field (CDFS), and the XMM Large Scale Structure Survey field (XMM-LSS). The XMM-LSS consists of two separate but spatially adjacent fields: the Subaru/XMM-Newton DEEP Survey field (XMM-SXDSS³) and the Canadian-France-Hawaii Telescope Legacy Survey field (XMM-CFHTLS⁴). In addition to the SWIRE fields we also include the COSMOS⁵ field for a total of five fields.

In all of the PRIMUS target fields we have near-UV (NUV) and far-UV (FUV) measurements from the *GALEX* Deep Imaging Survey (DIS; Martin et al. (2005); Morrissey et al. (2005)). To minimize contamination from neighboring sources, we use a Bayesian photometric code EM_{PHOT} (based on expectation maximization algorithm of Guillaume et al. (2006)). Furthermore, we use ground-based optical and *Spitzer*/IRAC mid-infrared photometric catalogs in each of the fields to obtain integrated fluxes in all photometric bands. To summarize,

¹ <http://swire.ipac.caltech.edu/swire/swire.html>

² <http://dipastro.pd.astro.it/esis>

³ <http://www.naoj.org/cience/SubaruProject/SDS>

⁴ <http://www.cfht.hawaii.edu/Science/CFHLS>

⁵ <http://cosmos.astro.caltech.edu>

TABLE 1
GALAXY SUBSAMPLES

Survey	Redshift (z)	Density	Quiescent	Star-forming
SDSS	0.0375 – 0.145	High	5470	4501
		Mid	3614	4438
		Low	5419	8927
PRIMUS	0.2 – 0.4	High	322	583
		Mid	177	403
		Low	768	2516
PRIMUS	0.4 – 0.6	High	350	675
		Mid	195	405
		Low	871	2385
PRIMUS	0.6 – 0.8	High	347	430
		Mid	186	327
		Low	833	1847
PRIMUS	0.8 – 1.0	High	136	232
		Mid	94	163
		Low	373	810

the general strategy employed is to use a circular aperture photometry to constrain the shape of the SED and then fixing the overall normalization to a estimate of the total magnitude in the detection band. Moustakas et al. (2013b) provides a detailed description of the calculation for each of our target fields.

From the spectroscopic redshift and photometry described above, we use *iSEDfit* to determine stellar masses, star formation rates (SFRs) and other physical properties in a simplified Bayesian framework. *iSEDfit*, which we will only briefly mention in this paper is discussed in detail in Appendix A. of Moustakas et al. (2013b). The code uses the redshift and the observed photometry of the galaxies to determine the statistical likelihood of a large ensemble of generated model SEDs. These generated model SEDs depend on population synthesis models and prior parameters. In order to derive our fiducial stellar masses and star formation rates, we use the Flexible Stellar Population Synthesis (FSPS) models (Conroy & Gunn (2010)) based on the Chabrier (2003) IMF. Other prior parameters are listed in Section 4.1 of Moustakas et al. (2013b). The photometric bands we use for the fitting in our PRIMUS data are the *GALEX* FUV and NUV, the two shortest IRAC bands at 3.6 and 4.5 μm , and the five optical bands (in the COSMOS field, we fit seven optical bands and near-infrared bands, see Section 2.3 of Moustakas et al. (2013b)). For a more detailed description of the data used in this paper, we refer readers to Moustakas et al. (2013b) Section 2 and Section 4.1.

2.2. SDSS-GALEX

For galaxies at low redshifts we use the SDSS Data Release 7 (DR7; Abazajian et al. (2009)). From the SDSS DR7 data, which provides high fidelity *ugriz* photometry and spectroscopic redshift, we specifically use galaxies from the New York University Value-Added Galaxy Catalog that satisfy the main sample criterion

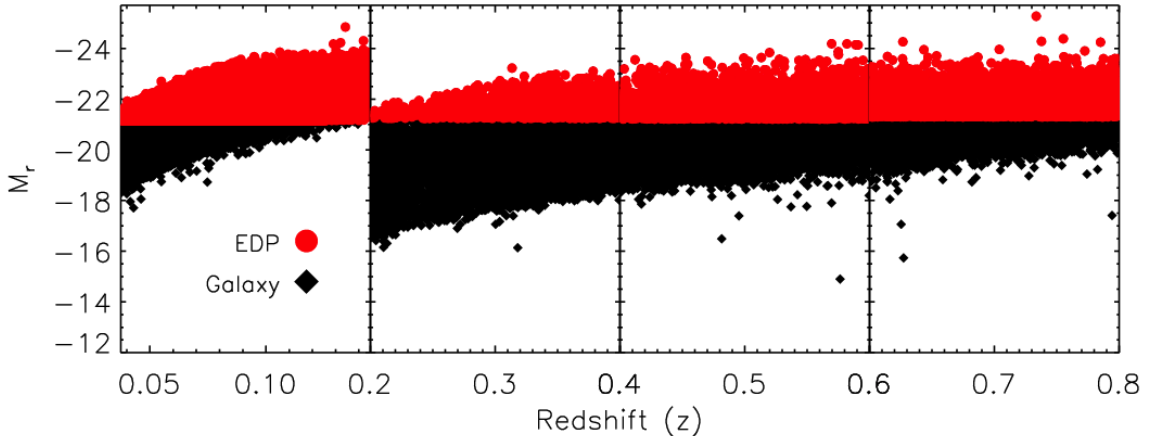


FIG. 1.— Absolute magnitude M_r versus redshift for the target galaxy population (black) with the Environment Defining Population (red) plotted on top. Both samples are divided into redshift bins: 0.0375–0.145, 0.2–0.4, 0.4–0.6, and 0.6–0.8. Stellar mass completeness limits are imposed on the target galaxy population (Section 2.3) while the M_r limits are imposed on the EDP (Section 2.5).

and have galaxy extinction corrected Petrosian magnitudes $14.5 < r < 17.6$ and spectroscopic redshifts within $0.01 < z < 0.2$ (Blanton et al. (2005b)). Within this sample, we select galaxies with medium depth observations from GALEX. This is done by first retrieving the positions of all GALEX tiles from GALEX Release 6 with total exposure time greater than 1 ks and then constructing a joint angular selection function of the SDSS-GALEX Sample. This results in a final sample of 167,727 SDSS galaxies with GALEX imaging.

For this final sample, we use the MAST/CasJobs⁶ interface and a 4'' diameter search radius to obtain the NUV and FUV photometry. For the optical photometry we use *ugriz* bands (from the SDSS model magnitudes) scaled to the *r*-band *cmodel* magnitude. We supplement the above UV and optical photometry with integrated *JHK_s* magnitudes from the 2MASS Extended Source Catalog (XSC; Jarrett et al. (2000)) and with photometry at 3.4 and 4.6 μm from the WISE All-Sky Data Release⁷. Further details on the photometry of the SDSS data used in this paper is provided in Moustakas et al. (2013b) Section 2.4.

To obtain the stellar masses and SFRs for the SDSS data we use *iSEDfit*, as we did for the PRIMUS data in Section 2.1. However, we use twelve photometric bands: GALEX FUV and NUV, SDSS *ugriz*, 2MASS *JHK_s*, and WISE 3.4 and 4.6 μm .

2.3. Target Galaxy Samples

In this section we will define the target galaxy sample we use to compute the SMFs and QFs. The population at $z \sim 0.5$ is derived from the PRIMUS data and the population at $z \sim 0.1$ is derived from the SDSS data, both described above. For the intermediate redshift target galaxy population, we begin with the selection criteria imposed in Moustakas et al. (2013b) for its parent sample. We take the statistically complete *primary* sample from the PRIMUS data (Coil et al. (2011)) and then

impose magnitude limits on optical selection bands as specified in Moustakas et al. (2013b) Table 1. These limits are in different optical selection bands and have different values for our five PRIMUS target fields. Next, we join the PRIMUS, GALEX, and IRAC window functions to construct the angular selection function. We then exclude stars and broad-line AGN to only select objects spectroscopically classified as galaxies, with high-quality spectroscopic redshifts ($Q \geq 3$) within the redshift range 0.2–0.8. Afterwards we assign statistical weights (described in Coil et al. (2011) and Cool et al. (2013)) to the objects in order to correct for targeting incompleteness and redshift failures. The statistical weight, w_i , for each galaxy is given by

$$w_i = (f_{\text{target}} \times f_{\text{collision}} \times f_{\text{success}})^{-1}, \quad (1)$$

Equation (1) in Moustakas et al. (2013b). In addition to the selection criteria above, we impose stellar mass limits to have a stellar mass complete sample. Stellar mass completeness limits for a magnitude-limited survey such as PRIMUS is a function of redshift, apparent magnitude limit of the survey, and the typical stellar mass-to-light ratio of galaxies near the flux limit.

As done in Moustakas et al. (2013b), we follow Pozzetti et al. (2010) to empirically determine the stellar mass completeness limits. Briefly, for each of the target galaxies we compute \mathcal{M}_{lim} using $\log \mathcal{M}_{\text{lim}} = \log \mathcal{M} + 0.4(m - m_{\text{lim}})$. \mathcal{M} is the stellar mass of the galaxy in units of \mathcal{M}_{\odot} , \mathcal{M}_{lim} is the stellar mass of each galaxy if its magnitude was equal to the survey magnitude limit, m is observed apparent magnitude in the selection band, and m_{lim} is the magnitude limit for our five fields mentioned above. We construct a cumulative distribution of \mathcal{M}_{lim} for the 15% faintest galaxies in $\Delta z = 0.04$ redshift bins. In each of these redshift bins, we calculate the minimum stellar mass that includes 95% of the galaxies. Separately for quiescent and star-forming galaxies, we fit quadratic polynomials to the minimum stellar masses versus redshift (classification scheme for quiescent and star-forming galaxies is included in the following section). Finally, we use the polynomials to obtain the minimum

⁶ <http://galex.stsci.edu/casjobs>

⁷ <http://wise2.ipac.caltech.edu/docs/release/allsky>

stellar masses at the center of the redshift bins, $0.2 - 0.4$, $0.4 - 0.6$, $0.6 - 0.8$, and $0.8 - 1.0$. These redshift bins will later be used to divide the target sample in order to compute the SMFs and QFs.

For our target galaxy population at low redshift we use our SDSS-*GALEX* data. We limit our target population within the redshift range $0.0375 - 0.145$ due to redshift limits later imposed for the volume-limited Environment Defining Population (Section 2.5). The NYU-VAGC provides statistical weight estimates for targeting completeness. In order to have a stellar mass complete low redshift target population, we use a uniform stellar mass limit of $10^9 \mathcal{M}_\odot$, which is above the surface brightness and stellar mass-to-light ratio completeness limits (Blanton et al. (2005a); Baldry et al. (2008)). The absolute magnitude (M_r) versus redshift for the target galaxy population at both low redshift and intermediate redshift is plotted in Figure 2.1. For further details we refer readers to Moustakas et al. (2013b) for details; more specifically, Figure 2., Sections 3.1, and 4.3.

2.4. Classifying Quiescent and Star-Forming Galaxies

With the galaxy sample defined in the previous section, we now classify the galaxies as quiescent or star-forming using an evolving cut based on specific star-formation rate utilized in Moustakas et al. (2013b) Section 3.2. To summarize, this classification method utilizes the star-forming (SF) sequence, which is the correlation between star-formation rate (SFR) and stellar mass in star-forming galaxies observed at least until $z \sim 2$. The PRIMUS sample displays a well-defined SF sequence within the redshift range of our target population. Using the power-law slope for the SF sequence derived by Salim et al. (2007) ($\text{SFR} \propto \mathcal{M}^{0.65}$) and the minimum of the quiescent/star-forming bimodality, determined empirically, we obtain the following equation to classify the target galaxies (Equation 2 in Moustakas et al. (2013b)):

$$\log(\text{SFR}_{\min}) = -0.49 + 0.64 \log(\mathcal{M} - 10) + 1.07(z - 0.1), \quad (2)$$

where \mathcal{M} is the stellar mass of the galaxy. If the target galaxy SFR and stellar mass place the galaxy above Equation 2 we classify it as star-forming; if below, as quiescent (Moustakas et al. (2013b) Figure 1.).

2.5. Galaxy Environment

In addition to the quiescent/star-forming classification, we now measure the environment for the target galaxy sample. First, we define the environment of a galaxy as the number of neighboring galaxies contained within a fixed aperture centered around it. As our aperture, we use a cylinder with dimensions with $R_{\text{ap}} = 2$ Mpc/h and $H_{\text{ap}} = 25$ Mpc/h. Though spherical apertures are often used in literature (e.g. Croton et al. (2005)), we use a cylindrical aperture in order to account for the PRIMUS redshift errors and redshift space distortions. Furthermore, we use a cylinder with the specified dimensions based on scale dependence for clustering analysis done in Blanton et al. (2006), which suggests that galactic properties such as star-formation histories are dependent on properties such as the masses of their host dark matter halos. Hence, with cylindrical apertures of $R_{\text{ap}} = 2$ Mpc/h we are able to probe the local small scale environment within the scales of the host dark matter

halos. We extend our analysis to cylindrical apertures with cylindrical apertures with $R_{\text{ap}} = 1$ Mpc/h; however we find that the variation in R_{ap} does not change our results significantly (difference in $f_Q < 0.05$). Similar fixed aperture methods have also been used in Croton et al. (2005) for galaxies in the 2dF Galaxy Redshift Survey (Colless et al. (2003)) and in Muldrew et al. (2012) for a mock galaxy catalogue generated from embedding galaxies onto the Millenium Dark Matter Simulation (Springel et al. (2005)).

In order to measure environment using an aperture defined above, we first construct a volume limited *Environment Defining Population* (EDP) with absolute magnitude cut-offs (M_r) at each redshift bin. The M_r cut-offs are selected so that the cumulative number density over M_r at all redshift bins are equal. By equating the cumulative number densities at different redshifts not only can we reasonably compare the environment values at different redshift bins but we also attempt to build an EDP that has a similar galaxy population in all of the redshift bins i.e. accounts for the progenitor bias. As Behroozi et al. (2013) and Leja et al. (2013) find in their analysis of the cumulative number density method using a dark matter simulation with abundance matching and a semi-analytic model, respectively, the cumulative number density method, while it does not precisely account for the scatter in mass accretion or galaxy-galaxy mergers, provides a reasonable means to compare galaxy populations over a wide range of cosmic time. Furthermore Behroozi et al. (2013)

To construct the EDP for the PRIMUS (hereafter PRIMUS EDP) we use the same PRIMUS data as the target galaxies, described in Section 2.3. We restrict the PRIMUS galaxies to $0.2 < z < 0.8$ and divide them into bins of $\Delta z = 0.2$. Before we consider the cumulative number densities in these bins, we first determine M_r limit for the highest redshift bin ($0.6 - 0.8$) by examining the M_r distribution with bin size $\Delta M_r = 0.25$ and select $M_{r,\text{lim}}$ near the peak of the distribution where bins with $M_r > M_{r,\text{lim}}$ have fewer galaxies than the bin at $M_{r,\text{lim}}$. We choose $M_{r,\text{lim}}(0.6 < z < 0.8)$ to be, conservatively, $M_r = -20.75$. For the other redshift bins, we impose absolute magnitude limits ($M_{r,\text{lim}}$) such that the cumulative number density of the bin ordered by M_r is equal to the cumulative number density of the highest redshift bin. The cumulative number density calculations accounts for the statistical weights of the galaxies (Section 2.1 and 2.2).

For the EDP of the SDSS-*GALEX* sample (hereafter SDSS EDP), we do not use the parent data of the SDSS-*GALEX* target sample the geometry of the combined angular selection function of the SDSS VAGC and *GALEX*. Instead, since FUV, NUV values are not necessary for EDP, we extend the parent data of the SDSS EDP to the entire SDSS VAGC to include galaxies outside of the *GALEX* window function. Furthermore, we impose a redshift range of $0.0375 - 0.145$ on the SDSS EDP. This redshift range is due to the lack of faint galaxies at $z \sim 0.2$ and the lack of bright galaxies at $z \sim 0.01$ in the SDSS VAGC data. The lower bound for the redshift range was empirically determined by the bright limit and the upper bound by the faint limit of the M_r versus redshift distribution. The same fixed cumulative number density method, described above, is used on this SDSS

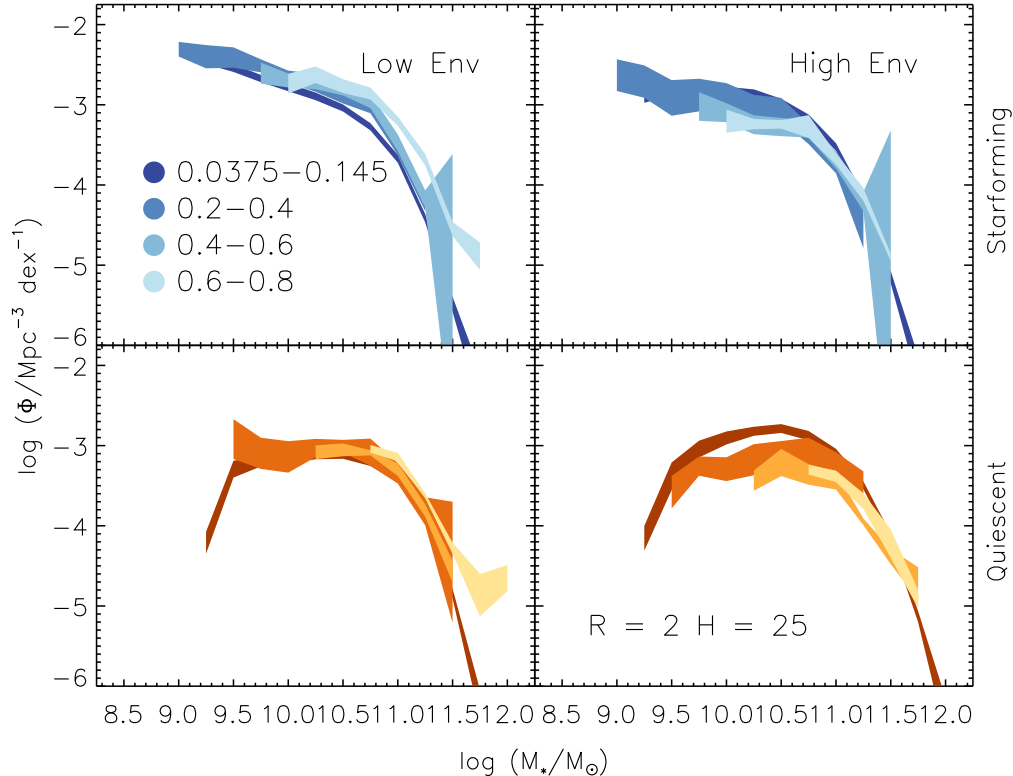


FIG. 2.— Evolution of stellar mass functions of star-forming (top) and quiescent (bottom) target galaxies in low (left) and high (right) environments from redshift range $z = 0 - 0.8$. The environment of each galaxy was calculated using a cylindrical aperture size of $R = 2\text{Mpc}$ and $H = 25\text{Mpc}$ and classification based on the cut-offs specified in Table 2. The SMFs use mass bins of width $\Delta\log(M/M_\odot) = 0.25$. In each panel we use shades of blue (star-forming) and orange (quiescent) to represent the SMF at different redshift, higher redshifts being progressively lighter.

EDP to equate the cumulative number density of SDSS EDP to the highest redshift bin of the PRIMUS EDP. Ultimately, we get $M_{r,\text{lim}} = -20.57, -20.73, -20.80$ and -20.95 for the redshift bins $0.06 - 0.145, 0.2 - 0.4, 0.4 - 0.6, 0.6 - 0.8$, respectively. These absolute magnitude limits are illustrated in Figure 2.1, which plots M_r distribution as a function of redshift for the EDP (red) and the target galaxy sample (black).

Finally with the EDP we measure the environment for each of the target sample galaxies by counting the number of EDP galaxies, n_{env} , with RA , Dec , and redshift within the aperture surrounding it. n_{env} accounts for the statistical weights of the EDP galaxies. More specific details for the dense and sparse environment cut-offs the various apertures are provided in Table 2.

2.6. Edge Effects

One of the challenges in obtaining the galaxy environment using a fixed aperture method is accounting for the edges of the survey. For galaxies located near the edge of the survey, part of the fixed aperture encompassing it will lie outside the survey regions. In this case, the n_{env} will only reflect the fraction of the environment within the survey geometry.

In order to account for these edge effects, we use a Monte Carlo method to impose edge cuts on the target galaxy population. We begin by computing the angular separation, θ_{ap} that corresponds to the radius of the

aperture at the redshifts of the target galaxies. Then the galaxies are matched to a sample of $N_{\text{ransack}} = 1,000,000$ points with RA and Dec randomly generated within the window function of the EDP. We refer to this randomly generated redshift-less sample as the “ransack” sample, based on the procedure used to construct them. For each target galaxy, we count the number of ransack points, n_{ransack} , within θ_{ap} of the galaxy’s RA and Dec value. The n_{ransack} values are then compared to the expected value:

$$E[n_{\text{ransack}}] = \frac{N_{\text{ransack}}}{A_{\text{EDP}}} \times \pi \theta_{\text{ap}}^2 \times f_{\text{thresh}} \quad (3)$$

where A_{EDP} is the total angular area of the target fields and f_{thresh} is the fractional threshold for the edge effect cut-off, which we vary based on R_{ap} (listed in Table 2). If n_{ransack} for a target galaxy is greater than $E[n_{\text{ransack}}]$ then the target galaxy remains in the sample; otherwise, it is discarded from the sample.

3. RESULTS

In this section, we describe our calculation to obtain the stellar mass functions (SMF) and quiescent fraction (QF) using the stellar mass complete target population. In Section .1 we construct the non-parametric estimate of SMFs for star-forming and quiescent galaxies in redshift bins: $0.06 - 0.145, 0.2 - 0.4, 0.4 - 0.6, 0.6 - 0.8$, and $0.8 - 1.0$. From these SMFs, we calculate the QFs in section 3.2.

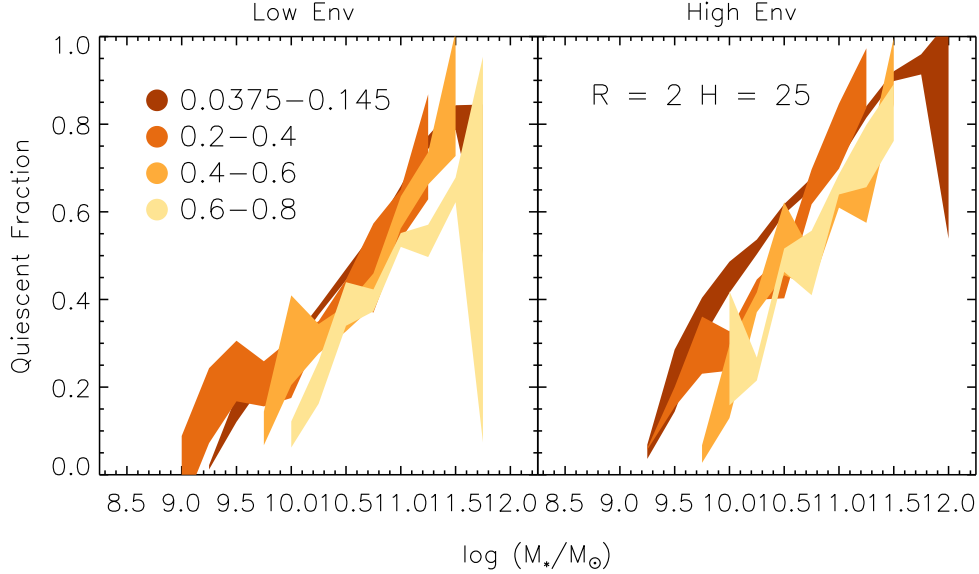


FIG. 3.— Evolution of the quiescent fraction for target galaxies in low (left) and (high) environments from redshift range $z = 0 - 0.8$. The QFs were calculated using the SMFs specified in Figure .1. We use lighter shades of orange for the QFs at higher redshifts.

3.1. Stellar Mass Function

Our primary interest in the SMFs is to investigate the evolution of the quiescent fraction in different density environments over our redshift range. Hence, we first divide the target galaxy sample into quiescent and star-forming target samples using the classification scheme from Section 2.4. Both the quiescent and star-forming target samples are then divided into “high-environment”, “mid-environment”, and “low-environment” based on their environment classification described in Section 2.5. We use the five redshift bins described in ?? to get a total of 30 subsamples that are classified based on star-forming/quiescent, environment, and redshift. For each of these subsamples, we calculate the SMFs and their errors.

To calculate the SMFs we employ a non-parametric $1/V_{\max}$ estimator commonly used for galaxy luminosity functions, as done in Moustakas et al. (2013b) and discussed in the review Johnston (2011). The differential SMF is given by the following equation:

$$\Phi(\log \mathcal{M})\Delta(\log \mathcal{M}) = \sum_{i=1}^N \frac{w_i}{V_{\max, \text{avail}, i}}. \quad (4)$$

The equation above is the same as Equation 3. in Moustakas et al. (2013b) other than the use of $V_{\max, \text{avail}}$ instead than V_{\max} , which is used because of edge effect considerations along the edges of the survey. w_i represents the statistical weight of each galaxy i and $\Phi(\log \mathcal{M})\Delta(\log \mathcal{M})$ is the number of galaxies (N) per unit volume within the stellar mass range $[\log \mathcal{M}, \log \mathcal{M} + \Delta(\log \mathcal{M})]$.

$V_{\max, i}$ is the maximum cosmological volume where it is possible to observe each galaxy i given the apparent magnitude limits of the survey. However in Section 2.6 we remove the galaxies that lie on the edge of from the target sample. In doing so we reduce the maximum cosmological volume where a galaxy can be observed, thereby reducing $V_{\max, i}$ to a what we will refer to as $V_{\max, \text{avail}, i}$.

To calculate $V_{\max, \text{avail}, i}$, we generate a sample of points

with random RA , Dec , and z (not to be confused with the ransack sample in Section 2.6) within the target fields and the redshift limits of the surveys. We then impose the same edge effect cut-off used on the target galaxy population, described in Section 2.6, onto this random sample. For each redshift bin j out of the n_{bin} redshift bins, we compute the fraction of the random points that remain in the bin after the points near the edges are removed: $f_{\text{edge}, j}$. Finally we calculate: $V_{\max, \text{avail}, i} = V_{\max, i} \times f_{\text{edge}, j}$. The $V_{\max, i}$ values in the equation above are computed following the method described in Moustakas et al. (2013b) with the same redshift-dependent K -correction from observed SED and luminosity evolution model. For more details on computing V_{\max} we refer readers to Section 4.2 in Moustakas et al. (2013b).

In order to calculate the uncertainty of the SMFs from the sample variance, we use a standard jackknife technique as done in Moustakas et al. (2013b). For the PRIMUS target galaxies, we calculate SMFs after excluding one of the target fields each time. Note in these SMF calculations, the $1/V_{\max, \text{avail}}$ is corrected for the fact that excluding a field changes the total area of the survey. Then using the calculated SMFs we calculate the uncertainty:

$$\sigma^j = \sqrt{\frac{M-1}{M} \sum_{k=1}^M (\Phi_k^j - \langle \Phi^j \rangle)^2} \quad (5)$$

M in this equation is the number of jack knife SMFs in the mass bin; this value is different for each of the stellar mass bins since not all mass bins have galaxies from all five target fields. $\langle \Phi^j \rangle$ is the mean number density of galaxies in each stellar mass bin for all of the jack knife Φ^j s.

Since the SDSS-*GALEX* data are contained in one field, we first divide the field into a 30×20 rectangular grid and then only keep the sectors with at least 100 galaxies in them. Then we use the same method

as above to compute the uncertainties through the jack knife technique. The panels in the top two rows of Figure .1 show the SMFs with uncertainties for all 30 subsamples using aperture with dimensions $R_{\text{ap}} = 1\text{Mpc}$ and $h_{\text{ap}} = 50\text{Mpc}$.

3.2. Quiescent Fraction

In the last section, we calculated the SMFs for all subsamples of the target galaxies classified by star-forming/quiescent, environment, and redshift. In this section, we use the Φ calculated for each subsample above to compute their respective QFs:

$$\text{QF} = \frac{\Phi_{\text{Q}}}{\Phi_{\text{SF}} + \Phi_{\text{Q}}}. \quad (6)$$

Φ_{Q} and Φ_{SF} are the total number of galaxies per unit volume in stellar mass bin $\Delta(\log \mathcal{M})$ (Equation 4) for the quiescent and star-forming subsamples, respectively. A value for the QF is not assigned in stellar mass bins where either the star-formation subsample or the quiescent subsample do not have any galaxies. The bottom panels of Figure .1 shows the QFs for the different redshift bins and environment density classifications.

In order to compare the QFs at different redshift for each of the environment classifications, we perform a linear least squares fit on the QF for all the QFs. Then we use the value of the linear fit at an empirically selected fiducial mass $\log \mathcal{M}/\mathcal{M}_{\odot} = 10.5$ (using fiducial mass of $\log \mathcal{M}/\mathcal{M}_{\odot} = 11.0$ does not notably change the out-

come). These values are used to highlight and attempt to quantify the differences of the QFs at the different redshifts, Figure .1

The figures and analysis described in this paper are done using galaxy environment determined from a fixed aperture with $r_{\text{ap}} = 1\text{Mpc}$ and $h_{\text{ap}} = 50\text{Mpc}$. The same analysis was repeated for various aperture dimensions: $r_{\text{ap}} = 0.5, 1, 2, 3\text{Mpc}$ and $h_{\text{ap}} = 25, 50\text{Mpc}$. Minor adjustments to the environment classification thresholds were adopted in these analyses for the smaller apertures ($r_{\text{ap}} = 0.5, 1\text{Mpc}$ and $r_{\text{ap}} = 25\text{Mpc}$). The results obtained from using these different are consistent with the results displayed in this paper.

4. DISCUSSION AND CONCLUSIONS

We have measured the SMFs and QFs using low redshift SDSS-*GALEX* galaxies and intermediate redshift PRIMUS galaxies. Specifically we analyzed the evolution of the QFs over the redshift range $0.0 - 1.0$ for galaxies in environment densities (Figure .1). We find that there is an expected increase in QF with decrease in the redshift for subsamples in all environment densities. More importantly we find that the change in QF over redshift is independent of the environment and remains relatively equal for all environments.

• Comparison to other works.

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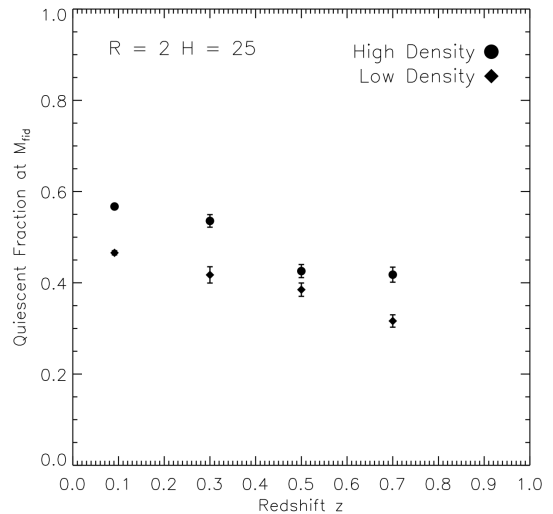


FIG. 4.— Quiescent fraction at fiducial mass $\log(\mathcal{M}/\mathcal{M}_{\odot}) = 10.5$ for low (square) and high (circle) environments in the redshift range $z = 0 - 0.8$. The evolution of $f_{\text{Q}}(\mathcal{M}_{\text{fid}})$ for galaxies in different environments shows an evolution over the redshift range that is independent of environment density. Furthermore throughout the redshift range the difference in $f_{\text{Q}}(\mathcal{M}_{\text{fid}})$ remains relatively constant and < 0.15 .

TABLE 2
FIXED CYLINDRICAL APERTURE DIMENSIONS

Radius (Mpc)	Height (Mpc)	n_{bin}	Edgecut	High Env Threshold (galaxies)	Low Env Threshold (galaxies)
1.0	50	6	80%	1.5	0.0
2.0	50	6	75%	4.0	0.0

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APPENDIX

Stellar Mass Function

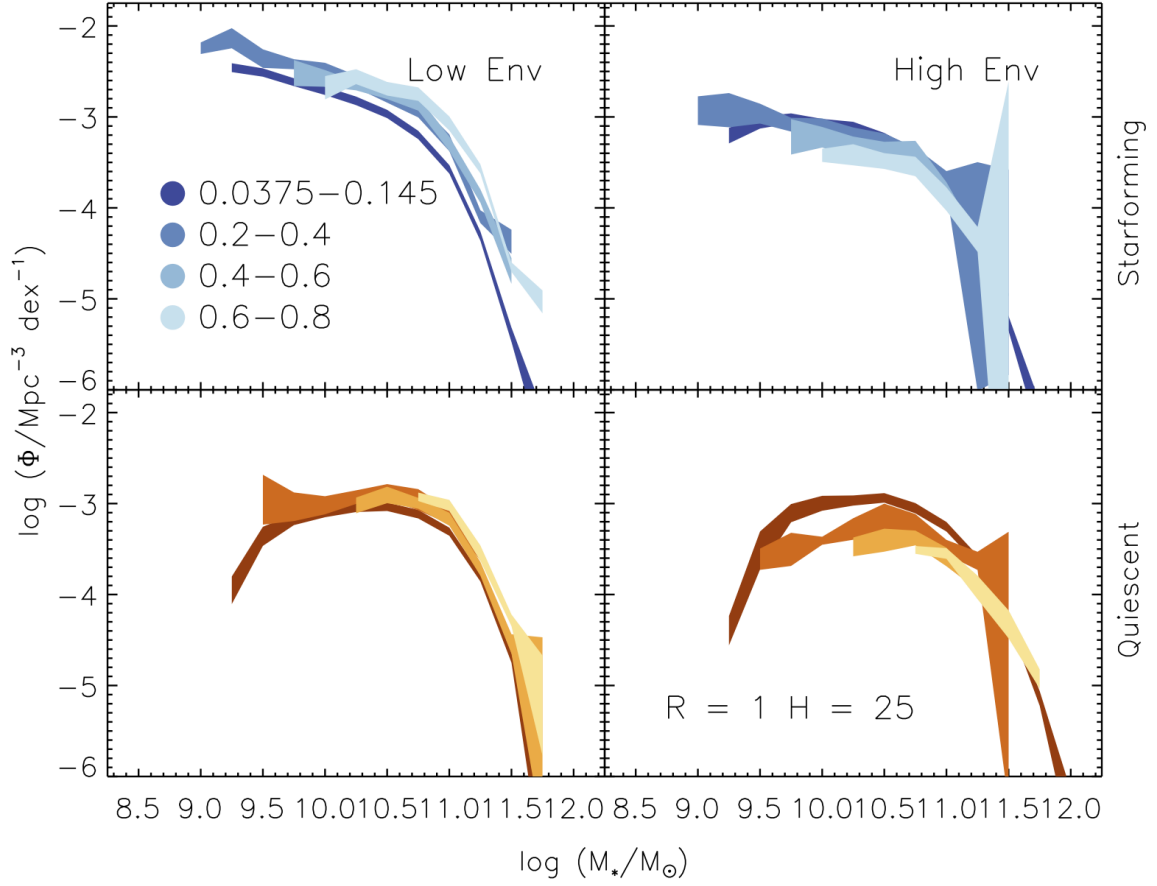


FIG. 5.— SMF for $r_{\text{ap}} = 1\text{Mpc}$ and $h_{\text{ap}} = 25\text{Mpc}$

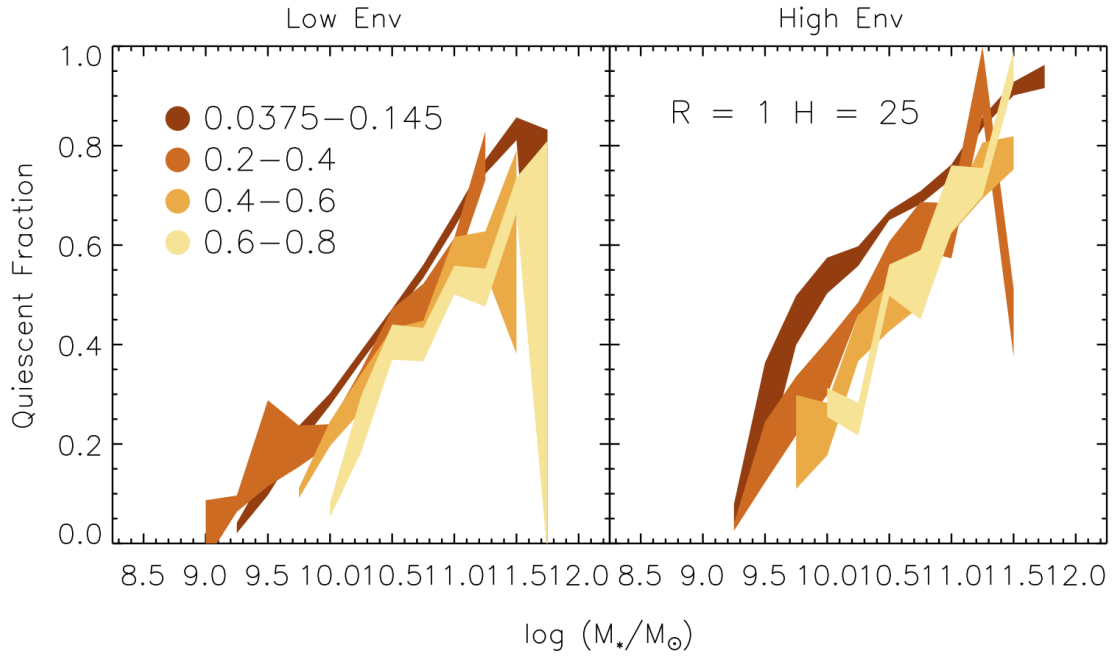


FIG. 6.— QF for $r_{\text{ap}} = 1\text{Mpc}$ and $h_{\text{ap}} = 25\text{Mpc}$

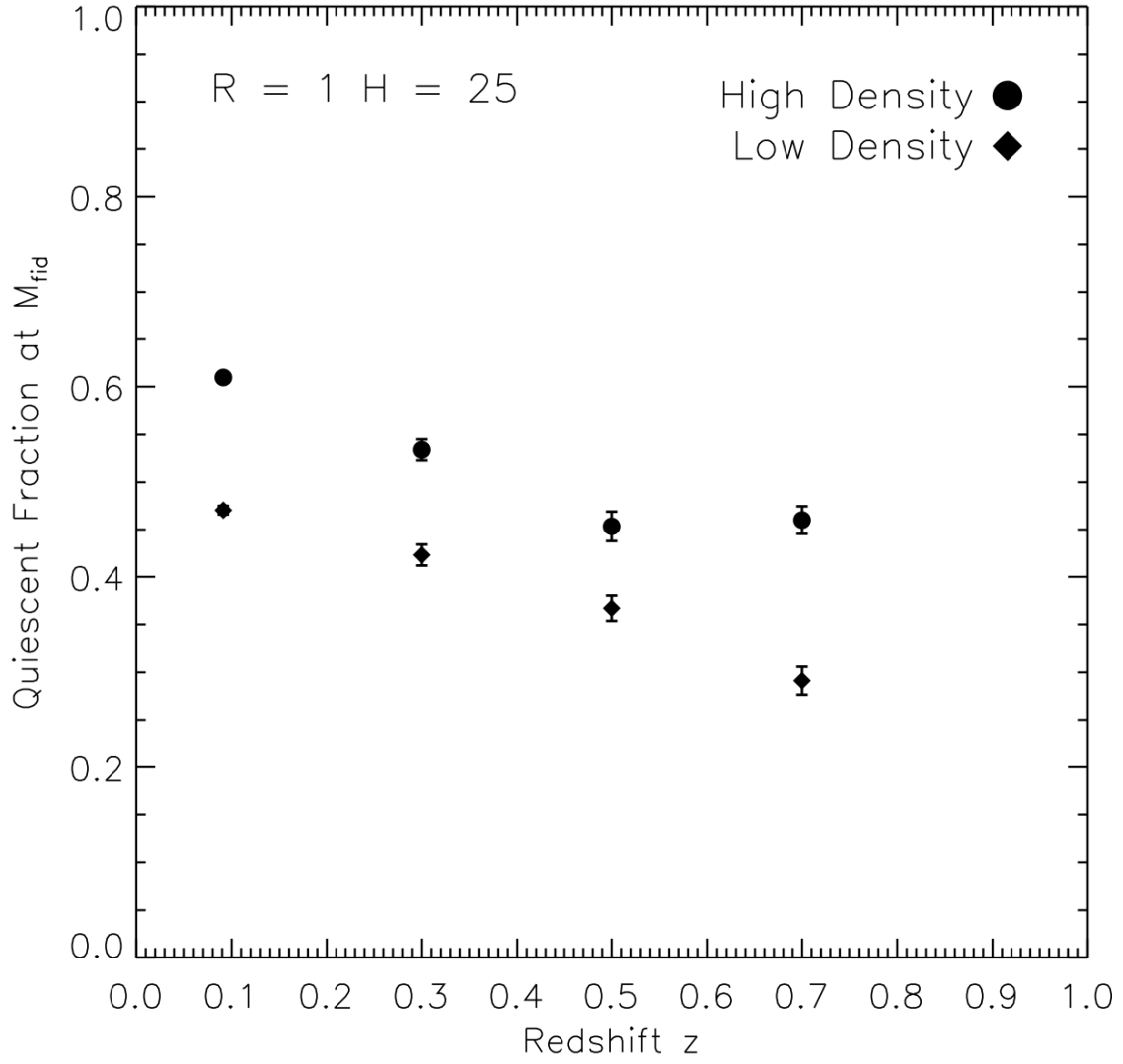


FIG. 7.— QF at fiducial mass for $r_{\text{ap}} = 2\text{Mpc}$ and $h_{\text{ap}} = 50\text{Mpc}$