

# PRIMUS: GALAXY ENVIRONMENT ON THE QUIESCENT FRACTION EVOLUTION AT $Z < 1$

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## ABSTRACT

We investigate the effects of galaxy environment on the evolution of the quiescent fraction ( $f_Q$ ) from  $z = 0.8$  to  $0.0$  using spectroscopic redshifts and multi-wavelength imaging data from PRIMUS Multi-object Survey (PRIMUS) and the Sloan Digital Sky Survey (SDSS). Our stellar mass limited galaxy sample consists of  $\sim 40000$  PRIMUS galaxies within  $z = 0.2 - 0.8$  and  $\sim 130000$  SDSS galaxies within  $z = 0.0375 - 0.145$ . We classify the galaxies as quiescent or star-forming, based on an evolving specific star formation cut, and as low or high density environments, based on fixed cylindrical aperture environment measurements on a volume-limited *Environment Definition Population* (from PRIMUS and SDSS).

For each of these subsamples, we examine the stellar mass function (SMF) evolution and compute the  $f_Q(\mathcal{M}_*)$ . We find that in both low and high density environments the quiescent fraction increases with cosmic time over the probed redshift range. Moreover, the difference between the quiescent fraction in low and high density environments remains constant throughout the redshift range. These results suggest that the evolution of the quiescent fraction is independent of environment and provide constraints on quenching mechanisms in high density environments such as groups and clusters.

## 1. INTRODUCTION

Galaxies, in their detailed properties, carry the imprints of their surroundings, with a strong dependence of the quiescent fraction of galaxies on their local environment (e.g. Hubble 1936; Oemler 1974; Dressler 1980; Hermit et al. 1996; Guzzo et al. 1997; for a recent review see Blanton & Moustakas 2009). The strength of this dependence is itself a strongly decreasing function of galaxy stellar mass; at the extreme, the lowest masses ( $< 10^9 M_\odot$ ) galaxies are quenched only in dense regions, and never in isolation (Geha et al. 2012). These effects vary with redshift at least in the densest clusters, as observed in the changing fraction of late-type spirals relative to the field found in studies of the morphology-density relation (Dressler 1984; Desai et al. 2007). Clearly understanding the properties of galaxies in the present-day universe requires a careful investigation of the role of environment, and how that role changes over time.

Nevertheless, the evolution of the role of environment is a relatively subtle effect and difficult to study. Although history of galaxies prior to  $z \sim 1$  appears to have been one of rapid assembly, since that time the galaxy population has continued to evolve, but less dramatically. Although there are detectable changes in the population, the major classes of galaxies existed at  $z \sim 1$ , in roughly the same relative numbers as today (Bundy et al. 2006; Borch et al. 2006; Taylor et al. 2009; Moustakas et al. 2013a). Furthermore, at those redshifts we can also de-

tect the dependence of galaxy properties on environment, with lower star-formation rate early-type galaxies populating the denser regions (Cooper et al. 2008; Patel et al. 2009; Kovač et al. 2010).

The most dramatic change in galaxy properties during the past eight billion years has been a remarkable decline in the star-formation rate of galaxies in the Universe (Hopkins & Beacom 2006). This decline appears dominated by decreases in the rates of star-formation of individual galaxies (Noeske et al. (2007)). There is evidence that a large fraction of the decline associated with strongly infrared-emitting starbursts (Bell et al. 2005; Magnelli et al. 2009). The decline does not appear to be due to the quenching of a large fractions of the star-forming population, as reflected in observations of the stellar mass function of quiescent and star-forming galaxies (Blanton et al. (2006), Bundy et al. 2006; Borch et al. 2006; Moustakas et al. 2013a). These findings leave little room for the participation of environmentally-driven quenching in the global census of star-formation. As Cooper et al. (2008) and others have pointed out, because the environmental dependence of total star-formation rates at fixed redshift is relatively small, environmentally effects are unlikely to cause the overall star-formation rate decline.

Thus, the impact of environment on galaxy formation has to be interpreted on top of the background of this overall decline affecting galaxies in all environments. The most straightforward investigation of would directly determine the star-forming properties of galaxies as a function of environment, stellar mass and redshift in a single, consistently analyzed data set. This analysis can reveal how galaxies are quenched in the universe over time, quantitatively establish the contribution of environmental effects to the overall trends, and reveal whether those trends happen equally in all environments. However, such an analysis has not been done previously due to

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the lack of sufficiently large samples. In this paper, we apply this approach using the Prism Multi-object Survey (PRIMUS; Coil et al. (2011), Cool et al. (2013)), the largest available redshift survey covering the epochs between  $0 < z < 1$ .

In Section 2 we present a brief description of the PRIMUS and SDSS data, our mass complete sample construction, and galaxy environment measurements. After dividing our galaxy sample into subsamples of star-forming or quiescent and high or low density environments, we compute and examine the evolution of the stellar mass functions for our subsamples. In Section 4, we calculate the quiescent fraction and compare the evolution of the quiescent fraction between high and low density environments. In Section 5 we discuss the trends observed in the stellar mass function and quiescent fraction evolution and discuss implications of the results on quenching mechanisms. Finally in Section 6 we summarize our results.

Throughout the paper we assume cosmology with  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$ , and  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

## 2. SAMPLE SELECTION

We are interested in quantifying the effects of galaxy environment on the evolution of the quenched population over the redshift range  $z < 1.0$ . For our analysis, we require a sample with sufficient depth and robustness to probe the redshift range and to measure galaxy environment. PRIMUS with its  $\sim 120,000$  spectroscopic redshifts provides ideal data at intermediate redshifts for our analysis. In addition, we anchor our analysis at intermediate redshifts with a low redshift sample derived from the Sloan Digital Sky Survey (?).

In Section 2.1 and Section 2.2 we provide a brief summary of the PRIMUS data and the SDSS data used for our sample selection. Then in Section 2.3 we define our stellar mass complete galaxy sample. Afterwards, in Section 2.4, we classify our sample galaxies as quiescent or star-forming then calculate the environment using a volume-limited *Environment Defining Population* in Section 2.5. Finally in Section 2.6, we correct our galaxy sample and environment measurements for edge effects of the surveys.

### 2.1. PRIMUS

At intermediate redshifts we use multiwavelength imaging and spectroscopic redshifts from PRIMUS, a faint galaxy survey with  $\sim 120,000$  precise spectroscopic redshifts ( $\sigma_z/(1+z) \approx 0.5\%$ ) within the range  $z \approx 0-1.2$ . The survey was conducted using a IMACS spectrograph on a Magellan I Baade 6.5 m telescope with a slitmask and low dispersion prism. For details on the PRIMUS observation methods such as survey design, targeting, and data summary, we refer readers to the survey papers: Coil et al. (2011) and Cool et al. (2013).

While the PRIMUS survey targeted seven distinct extragalactic deep fields for a total of  $\sim 9 \text{ deg}^2$ , we restrict our sample to five fields that have *GALEX* and *Spitzer*/IRAC imaging (similar to the sample selection in Moustakas et al. (2013b)). Four of these fields are a part of the *Spitzer* Wide-area Infrared Extragalactic Survey (SWIRE<sup>5</sup>): the European Large Area

TABLE 1  
GALAXY SUBSAMPLES

$z$	Environment	Quiescent	Star-forming
0.0375 – 0.145	High	5470	4501
	Low	5419	8927
0.2 – 0.4	High	322	583
	Low	768	2516
0.4 – 0.6	High	350	675
	Low	871	2385
0.6 – 0.8	High	347	430
	Low	833	1847

ISO Survey - South 1 field (ELAIS-S1<sup>6</sup>), the Chandra Deep Field South SWIRE field (CDFs), and the XMM Large Scale Structure Survey field (XMM-LSS). The XMM-LSS consists of two separate but spatially adjacent fields: the Subaru/XMM-Newton DEEP Survey field (XMM-SXDSS<sup>7</sup>) and the Canadian-France-Hawaii Telescope Legacy Survey field (XMM-CFHTLS<sup>8</sup>). Our fifth and final field is the Cosmic Evolution Survey (COSMOS<sup>9</sup>) field. For all of our fields we have near-UV (NUV) and far-UV (FUV) photometry from the *GALEX* Deep Imaging Survey (DIS; Martin et al. (2005); Morrissey et al. (2005)) as well as ground-based optical and *Spitzer*/IRAC mid-infrared photometric catalogs. Moustakas et al. (2013b) provides detailed descriptions of integrated flux calculations in the photometric bands for each of our fields.

Finally, from the spectroscopic redshift and broad wavelength photometry we use *iSEDfit*, a Bayesian SED modeling code, to calculate stellar masses and star formation rates (SFRs) for our sample galaxies. *iSEDfit* uses the redshift and the observed photometry of the galaxies to determine the statistical likelihood of a large ensemble of generated model SEDs. The model SEDs are generated using Flexible Stellar Population Synthesis (FSPS) models (Conroy & Gunn (2010)) based on Chabrier (2003) IMF, along with other prior parameters discussed in Section 4.1 and Appendix A. of Moustakas et al. (2013b). For the observed photometry, we use the *GALEX* FUV and NUV, the two shortest IRAC bands at 3.6 and 4.5  $\mu\text{m}$  (the two longer-wavelength IRAC channels are excluded because *iSEDfit* does not model hot dust or polycyclic aromatic hydrocarbons emission lines), and the optical bands.

### 2.2. SDSS-GALEX

At low redshifts, we use spectroscopic redshifts and high fidelity *ugriz* photometry from the SDSS Data Release 7 (DR7; Abazajian et al. (2009)). More specifically we select galaxies from the New York University Value-Added Galaxy Catalog (hereafter SDSS VAGC) that sat-

<sup>6</sup> <http://dipastro.pd.astro.it/esis>

<sup>7</sup> <http://www.naoj.org/science/SubaruProject/SDS>

<sup>8</sup> <http://www.cfht.hawaii.edu/Science/CFHLS>

<sup>9</sup> <http://cosmos.astro.caltech.edu>

<sup>5</sup> <http://swire.ipac.caltech.edu/swire/swire.html>



FIG. 1.— Absolute magnitude  $M_r$  versus redshift for the target galaxy population (black squares) with the Environment Defining Population (red circles) plotted on top. Both samples are divided into redshift bins:  $z \approx 0.0375 - 0.145$ ,  $0.2 - 0.4$ ,  $0.4 - 0.6$ , and  $0.6 - 0.8$  (panels left to right). Stellar mass completeness limits are imposed on the target galaxy population (Section 2.3) while the  $M_r$  limits are imposed on the EDP (Section 2.5). The lowest redshift bin ( $z \approx 0.0375 - 0.145$ ; leftmost panel) contain our galaxy sample and EDP selected from SDSS. The rest contain galaxies and EDP selected from PRIMUS.

isfy the main sample criterion and have galaxy extinction corrected Petrosian magnitudes  $14.5 < r < 17.6$  and spectroscopic redshifts within  $0.01 < z < 0.2$  (Blanton et al. (2005)). We further restrict the SDSS VAGC sample to only galaxies with medium depth observations with total exposure time greater than 1 ks from *GALEX* Release 6. This leaves 167,727 galaxies.

Next, we use the MAST/CasJobs<sup>10</sup> interface and a 4'' diameter search radius, to obtain the NUV and FNU photometry for the SDSS-*GALEX* galaxies. For optical photometry, we use the *ugriz* bands from the SDSS model magnitudes scaled to the *r*-band *cmodel* magnitude. These photometric bands are then supplemented with integrated *JHK<sub>s</sub>* magnitudes from the 2MASS Extended Source Catalog (XSC; Jarrett et al. (2000)) and with photometry at 3.4 and 4.6  $\mu$ m from the WISE All-Sky Data Release<sup>11</sup>. Further details regarding the SDSS-*GALEX* sample photometry can be found in Section 2.4 of Moustakas et al. (2013b). As previously done on the PRIMUS data in Section 2.1, we use *iSEDfit* to obtain the stellar masses and star formation rates for the SDSS-*GALEX* sample.

While the SDSS-*GALEX* data discussed above is derived from SDSS VAGC Data Release 7, since the release of the catalog, there have been significant improvements in the photometry. Notably, Blanton et al. (2011) find that the background subtractions in SDSS Data Release 7 has a size dependent bias in the galaxy fluxes, where galaxy fluxes are underestimated by  $\sim 1.5$  mag for galaxies with  $r_{50} \sim 100$  arcsec. In addition Bernardi et al. (2013) find that the luminosity, and ultimately stellar mass, is underestimated when they are derived from the SDSS *cmodel* magnitudes in comparison to PyMorph single-Sersic and Sersic-Exponential fits to the surface brightness profiles.

In order to quantify the effects of these photometric underestimations in our analysis, we replace our SDSS

fluxes in the *ugriz* band with *ugriz* fluxes from the NASA-Sloan Atlas (NSA) catalog, which incorporate the improved background subtraction presented in Blanton et al. (2011) and uses single-Sersic fit fluxes rather than the standard SDSS *cmodel* fluxes. Using the ratio of the luminosity derived from the improved photometry over the luminosity derived from the standard SDSS VAGC photometry, we impose a preliminary correction on the stellar mass values obtained from *iSEDfit* assuming a consistent mass-to-light ratio. This mass correction leads to a significant increase in the stellar mass function for  $\mathcal{M} > 10^{11} \mathcal{M}_{\odot}$  that matches the results from Bernardi et al. (2013). However, the effect of the mass correction was negligible for the quiescent fraction evolution results. As a result, we do not discuss the issues with photometric measurements any further in this paper. The effects of the discussed photometry methods on the luminosity function and stellar mass function will be examined in detail in future work.

### 2.3. Galaxy Sample

From the low redshift SDSS-*GALEX* and intermediate redshift PRIMUS data, in this section, we define our mass complete galaxy sample. We begin by imposing the parent sample selection criteria from Moustakas et al. (2013b). More specifically, we take the statistically complete *primary* sample from the PRIMUS data (Coil et al. (2011)) and impose magnitude limits on optical selection bands as specified in Moustakas et al. (2013b) Table 1. These limits are in different optical selection bands and have distinct values for the five PRIMUS target fields. We then exclude stars and broad-line AGN to only select objects spectroscopically classified as galaxies, with high-quality spectroscopic redshifts ( $Q \geq 3$ ). Lastly, we impose a redshift range of  $0.2 < z < 0.8$  for the PRIMUS galaxy sample, where  $z > 0.2$  is selected due to limitations from sample variance and  $z < 0.8$  is selected due to the lack of sufficient statistics in subsamples defined below.

For the PRIMUS objects that meet the above crite-

<sup>10</sup> <http://galex.stsci.edu/casjobs>

<sup>11</sup> <http://wise2.ipac.caltech.edu/docs/release/allsky>

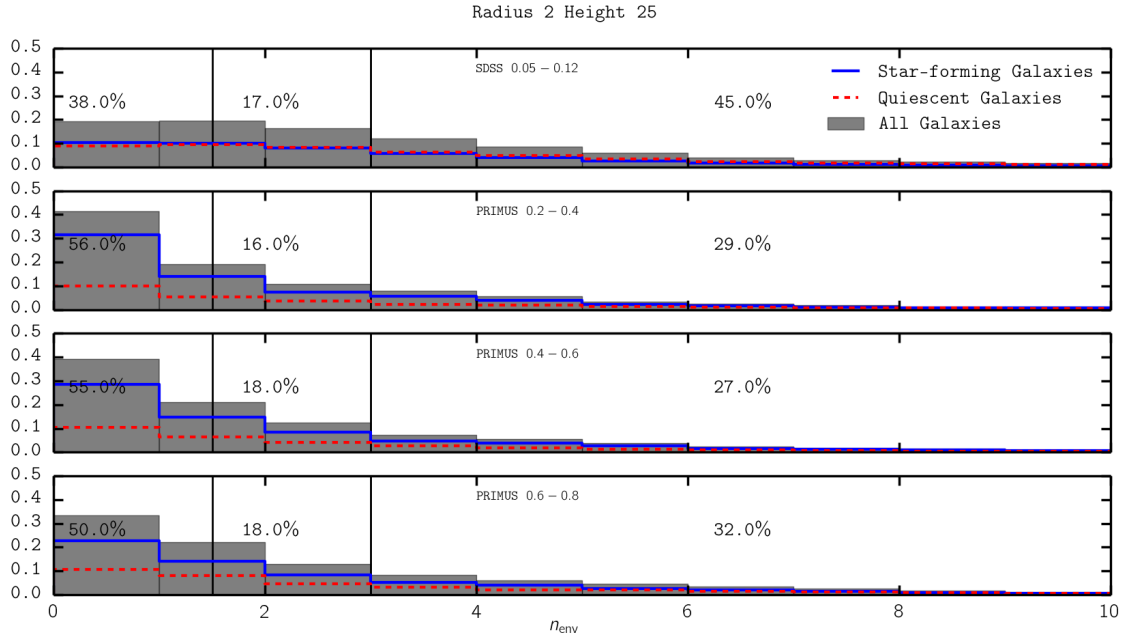


FIG. 2.— Environment measure distribution for our galaxy sample.

ria, we assign statistical weights (described in Coil et al. (2011) and Cool et al. (2013)) in order to correct for targeting incompleteness and redshift failures. The statistical weight,  $w_i$ , for each galaxy is given by

$$w_i = (f_{\text{target}} \times f_{\text{collision}} \times f_{\text{success}})^{-1}, \quad (1)$$

Equation (1) in Moustakas et al. (2013b). Since we are ultimately interested in a mass complete galaxy sample to derive SMFs and QFs, next we impose stellar mass completeness limits to our galaxy sample.

Stellar mass completeness limits for a magnitude-limited survey such as PRIMUS is a function of redshift, the apparent magnitude limit of the survey, and the typical stellar mass-to-light ratio of galaxies near the flux limit. As done in Moustakas et al. (2013b), we follow Pozzetti et al. (2010) to empirically determine the stellar mass completeness limits. For each of the target galaxies we compute  $\mathcal{M}_{\text{lim}}$  using  $\log \mathcal{M}_{\text{lim}} = \log \mathcal{M} + 0.4(m - m_{\text{lim}})$ , where  $\mathcal{M}$  is the stellar mass of the galaxy in  $\mathcal{M}_{\odot}$ ,  $\mathcal{M}_{\text{lim}}$  is the stellar mass of each galaxy if its magnitude was equal to the survey magnitude limit,  $m$  is the observed apparent magnitude in the selection band, and  $m_{\text{lim}}$  is the magnitude limit for our five fields. We construct a cumulative distribution of  $\mathcal{M}_{\text{lim}}$  for the 15% faintest galaxies in  $\Delta z = 0.04$  bins. In each of these redshift bins, we calculate the minimum stellar mass that includes 95% of the galaxies. Separately for quiescent and star-forming galaxies, we fit quadratic polynomials to the minimum stellar masses versus redshift (star-forming or quiescent classification is described in the following section). Finally, we use the polynomials to obtain the minimum stellar masses at the center of redshift bins, 0.2 – 0.4, 0.4 – 0.6, and 0.6 – 0.8, which are then used as PRIMUS stellar mass completeness limits.

For the low redshift portion of our galaxy sample, we start by limiting the SDSS-*GALEX* data to objects

within  $0.05 < z < 0.12$ , a redshift range later imposed on the volume-limited Environment Defining Population (Section 2.5). To account for the targeting incompleteness of the SDSS-*GALEX* sample, we use the statistical weight estimates provided by the NYU-VAGC catalog. Furthermore, we determine a uniform stellar mass completeness limit of  $10^{10.2} \mathcal{M}_{\odot}$  above the stellar mass-to-light ratio of the SDSS-*GALEX* data within the imposed redshift limits. We then apply this mass limit in order to obtain our mass-complete galaxy sample at low redshift.

We now have a stellar mass complete sample derived from SDSS-*GALEX* and PRIMUS data. Since our sample is derived from two different surveys, we account for the disparity in the redshift uncertainty. While PRIMUS provides precise spectroscopic redshifts with  $\sigma_z/(1+z) \approx 0.5\%$ , due to the higher redshift ranges it probes the uncertainties are significantly greater than that of the SDSS-*GALEX* redshifts. In order to have comparable environment measures throughout our redshift range, we apply PRIMUS redshift uncertainties to our galaxy sample selected from SDSS-*GALEX*. For each SDSS-*GALEX* galaxy, we adjust its redshift by randomly sampling a normal distribution with standard deviation  $\sigma = 0.005(1 + z_{\text{SDSS-GALEX}})$ , where  $z_{\text{SDSS-GALEX}}$  is the redshift of the galaxy.

We present the absolute magnitude ( $M_r$ ) versus redshift for the galaxy sample (black squares) in Figure 1. The left-most panel corresponds to the sample derived from the SDSS-*GALEX* data and the rest correspond to the target sample derived from the PRIMUS data divided in bins with  $\Delta z \sim 0.2$ .

#### 2.4. Classifying Quiescent and Star-Forming Galaxies

We now classify our mass complete galaxy sample into quiescent or star-forming using an evolving cut based on specific star-formation rate utilized in Moustakas et al. (2013b) Section 3.2. This classification method utilizes

the star-forming (SF) sequence, which is the correlation between star-formation rate (SFR) and stellar mass in star-forming galaxies observed at least until  $z \sim 2$  (Noeske et al. (2007)). The PRIMUS sample displays a well-defined SF sequence within the redshift range of our galaxy sample. Then using the power-law slope for the SF sequence from Salim et al. (2007) ( $\text{SFR} \propto \mathcal{M}^{0.65}$ ) and the minimum of the quiescent/star-forming bimodality, determined empirically, we obtain the following equation to classify the target galaxies (Equation 2 in Moustakas et al. (2013b)):

$$\log(\text{SFR}_{\min}) = -0.49 + 0.64\log(\mathcal{M} - 10) + 1.07(z - 0.1), \quad (2)$$

where  $\mathcal{M}$  is the stellar mass of the galaxy. If the target galaxy SFR and stellar mass lies above Equation 2.4 we classify it as star-forming; if below, as quiescent (Moustakas et al. (2013b) Figure 1.).

### 2.5. Galaxy Environment

We define the environment of a galaxy as the number of neighboring *Environment Defining Population* galaxies (defined below) within a fixed aperture centered around it. We use a fixed aperture measurement of environment because, as Muldrew et al. (2012) finds in their comparison of different environment definition using simulations, it provides a better probe of the entire halo in comparison to other definitions of environment such as nearest neighbor, which provides a better tracer at inter-halo scales.

For our aperture, we use a cylinder of dimensions:  $R_{\text{ap}} = 2 \text{ Mpc/h}$  and  $H_{\text{ap}} = 25 \text{ Mpc/h}$ . Though spherical apertures are often used in literature (e.g. Croton et al. (2005)), we use a cylindrical aperture in order to account for the PRIMUS redshift errors and redshift space distortions (i.e. "Finger of God" effect). Cooper et al. (2005) finds that  $\pm 1000 \text{ km s}^{-1}$  optimally reduces the effects of redshift space distortions. The PRIMUS redshift uncertainty at  $z \sim 0.7$  corresponds to  $\sigma_z < 0.01$ , so our aperture height choice of  $25 \text{ Mpc/h}$  accounts for both of these effects. Our choice of cylinder radius was motivated by scale dependence analyses in literature (Blanton et al. (2006), Wilman et al. (2010), and Muldrew et al. (2012)), which suggest that galactic properties such as color and quenched fractions are dependent on halo-scale properties such as host dark matter halo mass. Wilman et al. (2010), which uses environment defined by annuli of different radii, find positive correlation for quenched fraction and color on scales  $< 1 \text{ Mpc}$  and anti-correlation on scales  $> 3 \text{ Mpc}$ . Our choice of  $2 \text{ Mpc/h}$  provides sufficient sample size of galaxies in dense environments, for robust statistics, while tracing galactic properties within the halo scale.

Before we measure the environment for our galaxy sample, we first construct a volume limited *Environment Defining Population* (EDP) with absolute magnitude cut-offs ( $M_r$ ) at redshift bin of  $\Delta z \sim 0.2$ . The  $M_r$  cut-offs for the EDP are selected such that the cumulative number density over  $M_r$  for all redshift bins are equal. In addition to providing a reasonable comparison of environment measurements for wide redshift range, it also seeks to construct an EDP that contains similar galaxy populations through the redshift range (i.e. accounts for the progenitor bias). As Behroozi et al.

(2013) and Leja et al. (2013) find in their analysis of the cumulative number density method, though it does not precisely account for the scatter in mass accretion or galaxy-galaxy mergers, it provides a reasonable means to compare galaxy populations over a wide range of cosmic time.

In constructing the EDP for the PRIMUS (hereafter PRIMUS EDP) we use the same PRIMUS data used to select our galaxy sample (described in Section 2.3). We again restrict the PRIMUS galaxies to  $0.2 < z < 0.8$  and divide them into bins of  $\Delta z = 0.2$ . Before we consider the cumulative number densities in the redshift bins, we first determine the  $M_r$  limit for the highest redshift bin ( $0.6 - 0.8$ ) by examining the  $M_r$  distribution with bin size  $\Delta M_r = 0.25$  and select  $M_{r,\text{lim}}$  near the peak of the distribution where bins with  $M_r > M_{r,\text{lim}}$  have fewer galaxies than the bin at  $M_{r,\text{lim}}$ . We conservatively choose  $M_{r,\text{lim}}(0.6 < z < 0.8)$  to be  $M_r = -20.75$ . Then for the lower redshift bins, we impose absolute magnitude limits ( $M_{r,\text{lim}}$ ) such that the cumulative number density, calculated with the galaxy statistical weights, of the bin ordered by  $M_r$  is equal to the cumulative number density of the highest redshift bin at  $M_{r,\text{lim}}(0.6 < z < 0.8) = -20.75$ .

For the EDP of the SDSS-*GALEX* galaxy sample (hereafter SDSS EDP), we do not use the SDSS-*GALEX* parent data which uses the geometry of the combined angular selection function of the SDSS VAGC and *GALEX* (Section 2.2). Instead, since FUV, NUV values are not necessary for EDP, we extend the parent data of the SDSS EDP to the entire SDSS VAGC, including galaxies outside of the *GALEX* window function. Furthermore, we impose a redshift range of  $0.05 - 0.12$  on the SDSS EDP. This redshift range was empirically determined to account for the lack of faint galaxies at  $z \sim 0.2$  and the lack of bright galaxies at  $z \sim 0.01$  in the SDSS VAGC. The lower bound was determined by the bright limit and the upper bound by the faint limit of the  $M_r$  versus redshift distribution. As with the PRIMUS redshift bins, we determine the SDSS EDP  $M_{r,\text{lim}}$  by matching the cumulative number density of the highest redshift bin. For redshift bins  $0.05 - 0.12$ ,  $0.2 - 0.4$ ,  $0.4 - 0.6$ ,  $0.6 - 0.8$  we get  $M_{r,\text{lim}} = -20.57$ ,  $-20.73$ ,  $-20.80$  and  $-20.95$ , respectively. These absolute magnitude limits are illustrated in Figure 1, which show clear  $M_r$  cutoffs in the  $M_r$  distribution versus redshift for the EDP (red circles) on top of the target galaxy sample (black).

For our SDSS-*GALEX* galaxy sample, in Section 2.3, we imposed PRIMUS redshift errors in order to establish a consistent measurement of environment throughout our redshift range. We appropriately impose equivalent redshift adjustments for the SDSS EDP. So for the SDSS EDP galaxies that are within the SDSS-*GALEX* sample, we adjust the redshift by an identical amount. For the rest, we apply the same redshift adjustment procedure described in Section 2.3 in order to obtain PRIMUS level redshift uncertainties.

Finally we measure the environment for each galaxy in our galaxy sample by counting the number of EDP galaxies,  $n_{\text{env}}$  with  $RA$ ,  $Dec$ , and  $z$  within our cylindrical aperture centered around it.  $n_{\text{env}}$  accounts for the statistical weights of the EDP galaxies. Once we obtain environment measurements for all the galaxies in our galaxy sample, we classify galaxies with  $n_{\text{env}} < 1$  to be

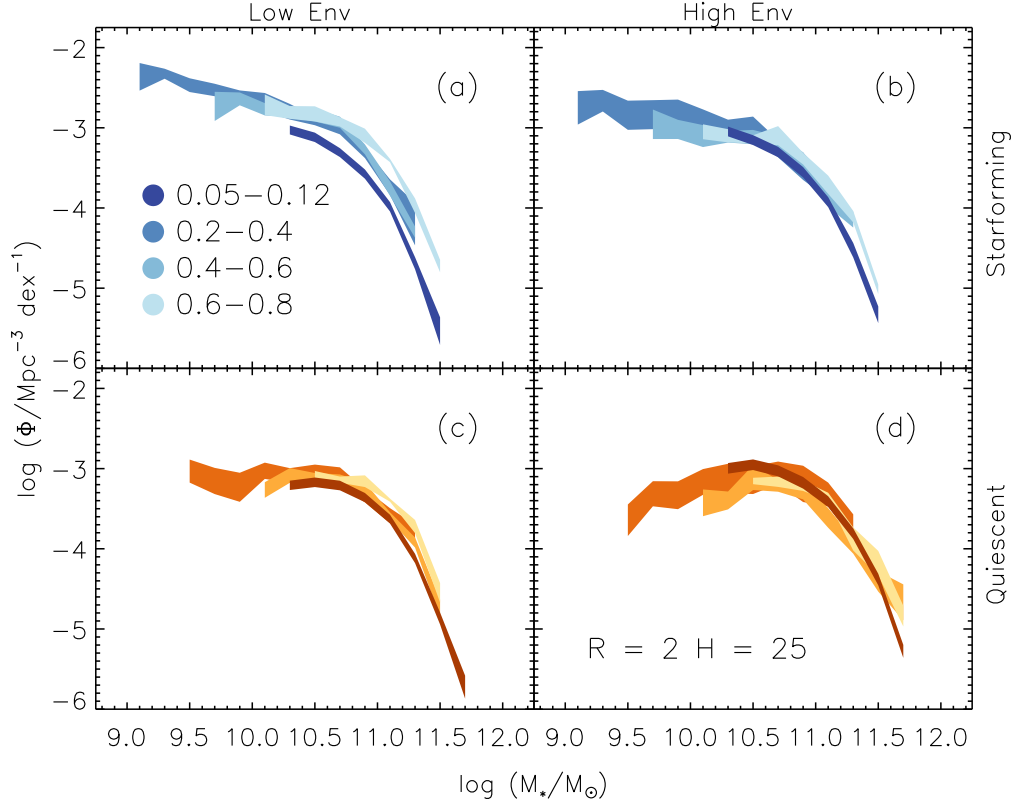


FIG. 3.— Evolution of stellar mass functions of star-forming (top) and quiescent (bottom) target galaxies in low (left) and high (right) environments from redshift range  $z = 0 - 0.8$ . The environment of each galaxy was calculated using a cylindrical aperture size of  $R = 2$  Mpc and  $H = 25$  Mpc and classification based on the cut-offs specified in Table 2. The SMFs use mass bins of width  $\Delta \log(M/M_\odot) = 0.25$ . In each panel we use shades of blue (star-forming) and orange (quiescent) to represent the SMF at different redshift, higher redshifts being progressively lighter.

”low” environment density and galaxies with  $n_{\text{env}} > 3$  to be ”high” environment densities. The high environment cutoff was selected in order to reduce contamination from galaxies in low environment densities while maintaining sufficient statistics.

The analysis we describe below uses a fixed cylindrical aperture with dimensions  $R_{\text{ap}} = 2$  Mpc and  $H_{\text{ap}} = 25$  Mpc to measure environment. The same analysis was extended for varying aperture dimensions  $R_{\text{ap}} = 1, 2, 3$  Mpc and  $H_{\text{ap}} = 25, 50$  Mpc. We also adjusted the environment classification thresholds for these analyses. The results obtained from using different apertures and environment classification thresholds are consistent with the results later presented.

### 2.6. Edge Effects

One of the challenges in obtaining accurate galaxy environments using a fixed aperture method is accounting for the edges of the survey. For galaxies located near the edge of the survey, part of the fixed aperture encompassing it will lie outside the survey regions. In this case,  $n_{\text{env}}$  only reflects the fraction of the environment within the survey geometry. To account for these edge effects, we use a Monte Carlo method to impose edge cutoffs on our galaxy sample. We begin by computing the angular separation,  $\theta_{\text{ap}}$  that corresponds to the radius of the fixed aperture (used to measure environment) at the redshifts of the galaxies. Then we compute the environment,

$n_{\text{ransack}}$ , of our galaxy sample with the EDP replaced by a ransack sample of  $N_{\text{ransack}} = 1,000,000$  points with  $RA$  and  $Dec$  randomly selected within the window function of the EDP. We compare the  $n_{\text{ransack}}$  values to the expected value computed based on the angular area of the environment defining aperture and the EDP window function:

$$E[n_{\text{ransack}}] = \frac{N_{\text{ransack}}}{A_{\text{EDP}}} \times \pi \theta_{\text{ap}}^2 \times f_{\text{thresh}}. \quad (3)$$

$A_{\text{EDP}}$  is the total angular area of the EDP window function and  $f_{\text{thresh}}$  is the fractional threshold for the edge effect cut-off, which we vary based on  $R_{\text{ap}}$  (listed in Table 2). If  $n_{\text{ransack}}$  for a target galaxy is greater than  $E[n_{\text{ransack}}]$  then the galaxy remains in our sample; otherwise, it is discarded.

## 3. STELLAR MASS FUNCTION

The target galaxy sample defined above has so far been classified into quiescent or star-forming and low or high density environments. We further divide these subsamples into redshift bins:  $0.0375 - 0.145$ ,  $0.2 - 0.4$ ,  $0.4 - 0.6$ , and  $0.6 - 0.8$  for a total 16 subsamples. In this section, we calculate the SMF for each of these subsamples.

To calculate the SMFs we employ a non-parametric  $1/V_{\text{max}}$  estimator commonly used for galaxy luminosity functions and stellar mass functions, as done in Moustakas et al. (2013b) and discussed in the review Johnston



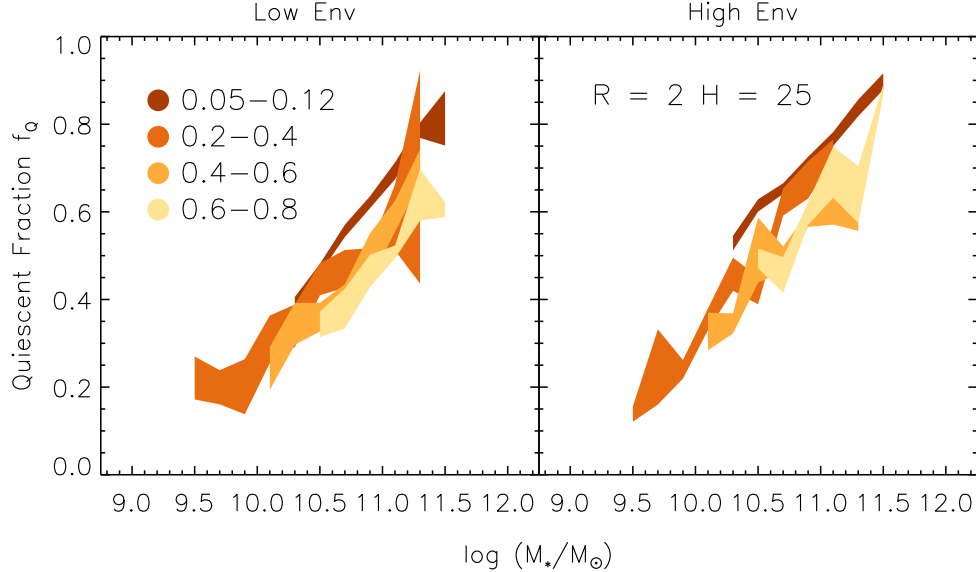


FIG. 4.— Evolution of the quiescent fraction  $f_Q$  for galaxies in low (left) and high (right) density environments for  $z < 0.8$ .  $f_Q$ s were calculated using the SMFs in Figure 3, as described in text. Darker shading indicates lower redshift.

(2011). The differential SMF is given by the following equation:

$$\Phi(\log \mathcal{M})\Delta(\log \mathcal{M}) = \sum_{i=1}^N \frac{w_i}{V_{\max, \text{avail}, i}}. \quad (4)$$

The equation above is same as Equation 3. in Moustakas et al. (2013b) except that we use  $V_{\max, \text{avail}}$  instead than  $V_{\max}$ , to account for the edge effects of the survey discussed in Section 2.6.  $w_i$  is the statistical weight of galaxy  $i$  and  $\Phi(\log \mathcal{M})\Delta(\log \mathcal{M})$  is the number of galaxies ( $N$ ) per unit volume within the stellar mass range  $[\log \mathcal{M}, \log \mathcal{M} + \Delta(\log \mathcal{M})]$ .

$V_{\max, i}$  is the maximum cosmological volume where it is possible to observe galaxy  $i$  given the apparent magnitude limits of the survey. However in Section 2.6 we remove galaxies that lie on the edge from our sample. In doing so we reduce the maximum cosmological volume where a galaxy can be observed, thereby reducing  $V_{\max, i}$ . We introduce the term  $V_{\max, \text{avail}, i}$  for the  $V_{\max, i}$  value that accounts for the survey edge effects.

To calculate  $V_{\max, \text{avail}, i}$ , we generate a sample of points with random  $RA$ ,  $Dec$  within the window function of our galaxy sample (SDSS-*GALEX* window function and the five PRIMUS fields) and random  $z$  within the redshift range. This is not to be confused with the *ransack* sample in Section 2.6. We then impose the same edge effect cutoffs we applied to our galaxy sample, using the same *ransack* environment calculation method, also discussed in Section 2.6. At redshift bins of  $\Delta z \sim 0.01$ , we compute the fraction of random points that remain in the bin after the edge effect cutoffs:  $f_{\text{edge}}$ . We then apply this factor to compute  $V_{\max, \text{avail}} = V_{\max} \times f_{\text{edge}}$ . The  $V_{\max}$  in the equation above are computed following the method described in Moustakas et al. (2013b) Section 4.2 with the same redshift-dependent  $K$ -correction from observed SED and luminosity evolution model.

In order to calculate the uncertainty of the SMFs from the sample variance, we use a standard jackknife tech-

nique (as done in Moustakas et al. (2013b)). For the PRIMUS galaxies, we calculate SMFs after excluding one of the five target fields at a time. For the SDSS target galaxies we divide the field into a  $12 \times 9$  rectangular  $RA$  and  $Dec$  grid and calculate the SMFs after excluding one grid each time. Using the calculated SMFs we calculate the uncertainty:

$$\sigma^j = \sqrt{\frac{M-1}{M} \sum_{k=1}^M (\Phi_k^j - \langle \Phi^j \rangle)^2} \quad (5)$$

$M$  in this equation is the number of jack knife SMFs in the stellar mass bin.  $\langle \Phi^j \rangle$  is the mean number density of galaxies in each stellar mass bin for all of the jack knife  $\Phi^j$ s.

In Figure 3, we present the SMFs for 16 target galaxy subsamples classified into quiescent/star-forming (orange/blue, bottom/top panels) and high/low environment density (left/right panels). The redshift evolution of the SMFs in each of these panels are indicated by a darker shade for lower redshift bins. The sample variance uncertainties are illustrated through the width of the SMFs.

Although we imposed PRIMUS redshift errors on our SDSS galaxies in order to consistently measure environment throughout our entire sample, we observe a significant discrepancy between the SDSS and PRIMUS samples in the fractional contribution of each environment classification within the redshift bin. For each of the PRIMUS redshift bins, approximately 50% of galaxies in the redshift bin are in low density environments and roughly 30% are in high density environments. In contrast, in the SDSS redshift bin, 38% of galaxies in the redshift bin are in low density environments and 45% are in high density environments. As a result, we caution that the SMF redshift evolution is inconsistent during the transition from the PRIMUS sample to the SDSS sample at  $z \sim 0.2$ .

Yet, from the PRIMUS samples alone, we probe a suf-

ficient redshift to observe notable trends in the SMF evolution for each of the panels in Figure 3. Later we compare these trends to SMF results from DEEP2 (Bundy et al. (2006)) and zCOSMOS (Bolzonella et al. (2010)). In panel (a), star-forming galaxies in low density environments, we find a significant decrease in the high mass end of the SMF ( $\mathcal{M} > 10^{10.75} \mathcal{M}_\odot$ ). In contrast at lower masses ( $\mathcal{M} < 10^{10.5} \mathcal{M}_\odot$ ) we observe a notable increase in the SMF. In panel (b), star-forming galaxies in high density environments, we notice a possible decrease above the knee of the SMF ( $\mathcal{M} \sim 10^{10.7} \mathcal{M}_\odot$ ) accompanied by an increase below the knee. For the quiescent population in low density environment, shown in panel (c) of Figure 3, we notice an increase in  $\Phi$  for lower masses, but we do not notice any clear SMF evolution at the massive end. Lastly for the quiescent population in high density environments, panel (d), we find significant increase in  $\Phi$  at all masses.

In comparison to Bundy et al. (2006) and Bolzonella et al. (2010), some of the noted trends are not immediately evident. However, this is mostly due to the conflicting mass limits proposed in each study. Not to mention, only part of the redshift range probed by Bundy et al. (2006) overlap with our results. Overall, we find that the trends we note in the SMF evolution above are in good agreement with results from DEEP2 and zCOSMOS.

Observing the evolutionary trends in SMF for each of these sub-populations simultaneously provides a naive narrative of the different galaxy evolutionary tracks involving environment and quenching. For example, the decrease in the massive star-forming galaxies in low density environment over cosmic time can be attributed to the transition of those galaxies to any of the other panels. The star-forming galaxies in low density environments that have quenched their star formation over time are possibly responsible for the increase of the quiescent, low density environment SMF over time. The star-forming galaxies that fall into higher density environments explain the increase in the star-forming high density environment SMF below the knee. Finally, star-forming galaxies in high density environments that have quenched their star-formation, quiescent galaxies that have transitioned from low to high density environments, and star-forming galaxies in low density environments that quench their star-formation while infalling to high density environments all contribute to the overall increase of the high environment quiescent SMF.

In addition to the evolution over cosmic time, we observe noticeable trends when we compare the SMFs for star-forming and quiescent galaxies between the two environments. Comparison of the SMFs in low versus high density environments reveal a noticeable relation between mass and density, with SMFs in high density environments having more massive galaxies. We further confirm this trend when we compare the median mass between the two environments to find that the median mass for galaxies in high density environments is significantly greater than in low density environments. The relationship between mass and environment observed in our SMFs mirror the well-established mass-density relation and observed mass segregation with environment in the literature (Bundy et al. (2006); Scodreggio et al. (2009); Bolzonella et al. (2010)).

We caution readers and reserve detailed analysis of the SMFs to future work due to the improvements in photometry that significantly alter the SDSS SMFs mentioned in Section 2.2.

#### 4. QUIESCENT FRACTION

The SMFs calculated in the previous section illustrate the stellar mass distribution of our galaxy population and its evolution over cosmic time. In this section, using the SMFs of our subsamples, we compare the quiescent and the star-forming populations by calculating the fraction of galaxies that have quenched their star-formation, the quiescent fraction. With the sufficient statistics available from SDSS and PRIMUS, we evaluate the quiescent fraction in bins of stellar mass and redshift for low and high density environments. More specifically, by analyzing the quiescent fraction with respect to mass, redshift, and environment we are able to compare the quiescent fraction evolution in low and high density environments. Through this comparison we reveal the subtle environmental effects in quenching star formation that are often obscured by the underlying correlations among galactic properties such as color-mass and mass-density relations (Cooper et al. (2010)). Furthermore by quantifying the environmental effects of star-formation quenching, we are able to provide constraints on the various proposed environmental quenching mechanisms, such as strangulation and ram-pressure stripping (McCarthy et al. (2008)).

From the SMF number densities ( $\Phi$ ) computed in the previous section, the quiescent fraction is computed as follows,

$$f_Q = \frac{\Phi_Q}{\Phi_{SF} + \Phi_Q}. \quad (6)$$

$\Phi_Q$  and  $\Phi_{SF}$  are the total number of galaxies per unit volume in stellar mass bin of  $\Delta(\log \mathcal{M}) = 0.25$  for the quiescent and star-forming subsamples, respectively (Equation 4). We compute  $f_Q$  for high and low density environments for all redshift bins as plotted in Figure 4, which shows the evolution of  $f_Q$  for high (right) and low (left) density environments. As in Figure 3, the darker shading represent lower redshift bins.

Most noticeably in Figure 4,  $f_Q$  increases monotonically as a function of mass at all redshift bin and environment. In other words, for all subsamples, galaxies in higher mass bins are more likely to have quenched its star-formation. With the roughly linear correlation between galaxy SFR to galaxy color and morphology, we find that this trend reflects the well established color-mass and morphology-mass relations: more massive galaxies are more likely to be red or early-type (Blanton & Moustakas (2009)).

Focusing on the redshift evolution of  $f_Q$ , we find that for both environments  $f_Q$  increases as redshift decreases. For high density environment, this is analogous to the Butcher-Oemler Effect (Butcher & Oemler (1984)), which states that galaxy populations in groups or clusters have higher  $f_{\text{blue}}$ , or lower  $f_Q$ , at higher redshift. For all environments a larger portion of our galaxy population has quenched its star-formation over time. Coupled with the decline of SFR in the star-forming population (Noeske et al. (2007); Cooper et al. (2008)), our results strongly agree with the observed decline of global star-formation in the universe (Hopkins & Bea-



com (2006)).

In addition, when we compare the stellar masses at which  $f_Q = 0.5$  for each subsample, we find that this mass decreases over cosmic time. This corresponds to the well-known mass-downsizing pattern in the literature (e.g. Bundy et al. (2006)). Furthermore, the mass-downsizing trend observed in each of our environment subsample is consistent with the trend observed in zCOSMOS Redshift Survey for isolated and group galaxies (Iovino et al. (2010)).

Finally, we compare our  $f_Q$  results to previous  $f_Q/f_{\text{red}}$  measurements for analogous environment classifications. For our lowest redshift bin, we find that our  $f_Q$  values for low and high environments are consistent with previous  $f_Q/f_{\text{red}}$  measurements for SDSS galaxies as a function of environment. For example, Tinker et al. (2011), using a group-finding algorithm on the SDSS DR7, presents the relationship between  $f_Q$  and overdensity for galaxies within the mass range  $\log \mathcal{M} = [9.8, 10.1]$ . The Tinker et al. (2011)  $f_Q$  at the lowest and highest overdensities,  $f_Q \sim 0.4$  and  $f_Q \sim 0.6$  respectively, are consistent with our  $f_Q$  for low and high density environment at the lowest mass bin ( $\log \mathcal{M} \sim 10.2$ ). At higher masses, our  $f_Q$  results are consistent with Baldry et al. (2006). Baldry et al. (2006) uses project neighbor density environment measures ( $\log \Sigma$ ) to obtain  $f_Q(\mathcal{M})$  for a range of environmental densities. Although the different environment measurements make a direct comparison unclear, at higher environments ( $\log \Sigma > 0.2$ ) in Baldry et al. (2006)  $f_Q \sim 0.6$  at  $\mathcal{M} \sim 10^{10.2} M_\odot$  and  $f_Q \sim 0.9$  at  $\mathcal{M} \sim 10^{11.5} M_\odot$ , in agreement our high density environment. Likewise, for lower environments in Baldry et al. (2006) ( $\log \Sigma < -0.4$ )  $f_Q \sim 0.4$  at  $\mathcal{M} \sim 10^{10.2} M_\odot$  and  $f_Q \sim 0.8$  at  $\mathcal{M} \sim 10^{11.5} M_\odot$ , which also agree with our low density environment  $f_Q$ .

For  $z > 0.2$ , we compare our PRIMUS  $f_Q$  results to the  $f_{\text{red}}$  (or  $1 - f_{\text{blue}}$ ) results from the zCOSMOS Redshift Survey (Iovino et al. (2010), Cucciati et al. (2010), Kovač et al. (2014)), which approximately covers the same redshift range as PRIMUS. Iovino et al. (2010), Cucciati et al. (2010), and Kovač et al. (2014) using a mass-complete galaxy sample derived from zCOSMOS and a group catalog, 3D local density contrast, and overdensity environment measurements, respectively, compare  $f_{\text{red}}$  with respect to environment. The  $f_{\text{blue}}$  for group and isolated galaxies from Iovino et al. (2010) are generally inconsistent with our  $1 - f_Q$  for high and low density environments. Similarly,  $f_{\text{red}}$  for high and low overdensities in Kovač et al. (2014) are greater overall than the PRIMUS  $f_Q$  values in high and low density environments. However, Kovač et al. (2014) points out that there is a significant difference between classifying the quenched population using color and SFR.  $f_Q$  defined by color is greater than  $f_Q$  defined by SFR by roughly 0.2. Although the propagation of this difference for different environments is not explored in Kovač et al. (2014), if we account for the difference evenly for  $f_{\text{red}}$  at all environments, the Kovač et al. (2014) results are generally consistent with our  $f_Q$  at high and low density environments.

Additionally, the zCOSMOS results indicate that the  $f_{\text{red}}$  values in high density environments are greater overall than in low density environments. The same trend

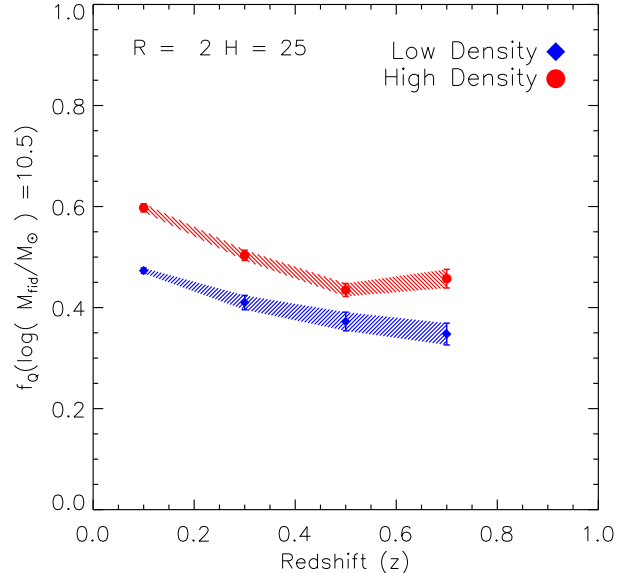


FIG. 5.— The evolution of the quiescent fraction at fiducial mass,  $f_Q(\mathcal{M}_{\text{fid}} = 10^{10.5} M_\odot)$ , for low (square) and high (circle) density environments within the redshift range  $z = 0.0 - 0.8$ . There is a significant increase in  $f_Q(\mathcal{M}_{\text{fid}})$  with decrease in redshift for both environments. In addition, over the entire redshift range, high density environment  $f_Q(\mathcal{M}_{\text{fid}})$  is greater than low density environment  $f_Q(\mathcal{M}_{\text{fid}})$ . However the difference in  $f_Q(\mathcal{M}_{\text{fid}})$  for the two environments remains constant throughout suggesting that while the quiescent fraction is higher in dense environments, the evolution of the quiescent fraction is independent of environment.

is observed in the PRIMUS results and will be discussed in detail later in the section. Another significant trend observed in the zCOSMOS results is the significant mass dependence of the  $f_{\text{red}}$  and  $f_{\text{blue}}$  evolution. Iovino et al. (2010) by examining  $f_{\text{blue}}(z)$  in different mass bins for group and isolated galaxies find significantly larger  $f_{\text{blue}}$  evolution at lower masses. In fact, Iovino et al. (2010) finds no evolution in  $f_{\text{blue}}$  for  $\mathcal{M} > 10^{10.9} M_\odot$  for both group and isolated galaxies over the redshift range  $0.4 < z < 0.9$ . On the other hand, in the lowest mass bin ( $10^{10.0} M_\odot < \mathcal{M} < 10^{10.5} M_\odot$ ) Iovino et al. (2010) finds  $f_{\text{blue}}(z \sim 0.4) - f_{\text{blue}}(z \sim 0.2) \sim 0.2$  for isolated galaxies and  $f_{\text{blue}}(z \sim 0.4) - f_{\text{blue}}(z \sim 0.2) \sim 0.1$  for group galaxies. In contrast, in each of the panels of Figure 4, we qualitatively find little mass dependence on the  $f_Q$  evolution for either low or high density environments. However, due to the distinct stellar mass completeness limits imposed at each redshift bin of our sample, subtle effects such as the mass dependence of the  $f_Q$  evolution are difficult to discern.

In order to more quantitatively compare the  $f_Q$  evolution, we fit  $f_Q$  for each subsample to a power-law parameterization as a function of stellar mass,

$$f_Q(\mathcal{M}_*) = a \log \left( \frac{\mathcal{M}_*}{10^{10.5} M_\odot} \right) + b, \quad (7)$$

where  $a$  and  $b$  are best-fit parameters using *MPFIT* (Markwardt (2009)). The value  $10^{10.5} M_\odot$  in the equation represents an empirically selected fiducial mass  $\mathcal{M}_{\text{fid}}$  within the stellar mass limits where there is a sufficiently large number of galaxies. This fiducial mass serves to

highlight and quantify the  $f_Q$  evolution for different redshifts. Figure 5 shows the evolution of  $f_Q(\mathcal{M}_{\text{fid}})$  from  $z \sim 0.7$  to  $\sim 0.1$  for low (diamond) and high (circle) density environments.

Figure 5 illustrates that throughout the entire redshift range, high density environment has a significantly greater  $f_Q(\mathcal{M}_{\text{fid}})$  than the lower density environment. This is in strong agreement with the zCOSMOS results from Cucciati et al. (2010) and Kovač et al. (2014). For  $z < 1.0$ , galaxy populations in higher density environments have higher quenched fractions. As galaxy color and SFR are related, our results support the existence of the color-density relation (Cucciati et al. (2010); Cooper et al. (2010)) in favor of the idea that the color-density relation is a merely a reflection of the mass-density or morphology-density relation (Scodreggio et al. (2009), Tasca et al. (2009)).

The same fitting procedure was performed using  $\mathcal{M}_{\text{fid}} = 10^{11} \mathcal{M}_{\odot}$ ; however, aside from an overall shift in  $f_Q(\mathcal{M}_{\text{fid}})$  by  $\sim 0.2$ , we observe the same evolutionary trends.  $f_Q(\mathcal{M}_{\text{fid}} = 10^{11} \mathcal{M}_{\odot})$  for both low and high density environments each increase by  $\sim 0.2$  as redshift decreases from  $\sim 0.7$  to  $\sim 0.1$ , just as Figure 5. Evidently, increasing the fiducial mass to  $10^{11} \mathcal{M}_{\odot}$  does not change the total evolution for either environments. On the other hand, Iovino et al. (2010), for both group and isolated galaxies, finds significantly less  $f_Q$  evolution for  $\mathcal{M} \sim 10^{11.0} \mathcal{M}_{\odot}$  than in  $\mathcal{M} \sim 10^{10.5} \mathcal{M}_{\odot}$ . Furthermore, Iovino et al. (2010) also finds that the difference in  $f_{\text{blue}}$  evolution between group and isolated galaxies ( $f_{\text{blue,isolated}}(z) - f_{\text{blue,group}}(z)$ ) is much greater at  $\mathcal{M} \sim 10^{10.5} \mathcal{M}_{\odot}$  than at  $\mathcal{M} \sim 10^{11} \mathcal{M}_{\odot}$ , where they find almost no difference. In contrast, we find that changing the fiducial mass does not significantly alter the difference in  $f_Q(\mathcal{M}_{\text{fid}})$ . The zCOSMOS results find a clear mass dependence in the  $f_{\text{blue}}$  evolution; however, the consistency in evolutionary trends when the fiducial mass is change reveal that our  $f_Q$  evolution exhibits little mass dependence. Although our sample from SDSS and PRIMUS provides significantly larger statistics, due to the limited mass ranges probed in our mass-complete sample (e.g.  $\mathcal{M} > 10^{10.5} \mathcal{M}_{\odot}$  for our  $z \sim 0.7$  bin) it is difficult to dismiss the mass dependence observed in the zCOSMOS results.

Finally, from Figure 5, we note that for both environments  $f_Q(\mathcal{M}_{\text{fid}})$  increases as redshift decreases. This highlights the general increase of  $f_Q(\mathcal{M})$  as redshift decreases noted earlier. More importantly, we quantify the total evolution of  $f_Q$  for both low and high density environments through  $f_Q(\mathcal{M}_{\text{fid}}, z \sim 0.1) - f_Q(\mathcal{M}_{\text{fid}}, z \sim 0.7)$  for both environments. Remarkably,  $f_Q$  evolves by the same amount for both environments. In other words, the total evolution of  $f_Q(\mathcal{M}_{\text{fid}})$  from  $z \sim 0.7$  to  $\sim 0.1$  show little environment dependence. This disagrees with the  $f_{\text{blue}}$  results from Iovino et al. (2010), which show a clear distinction between group and isolated galaxies. In Iovino et al. (2010),  $f_{\text{blue}}$  for isolated galaxies decrease much faster than  $f_{\text{blue}}$  for group galaxies. At  $\mathcal{M} \sim 10.5$ , Iovino et al. (2010) finds that  $f_{\text{blue}}$  for isolated galaxy decreases by 0.3 from  $z \sim 0.6$  to 0.3, while for  $f_{\text{blue}}$  for group galaxies only decrease by 0.1 in the same redshift limit, showing a clear dependence in environment. Considering that fixed aperture definition

of environment is susceptible to contamination due to PRIMUS redshift errors, we consider more stringent classification of high density environments extending the cut-off to  $n_{\text{env}} > 6, 9$ . While a more stringent classification increases the  $f_Q(\mathcal{M}_{\text{fid}})$  overall as well as the uncertainties, it does not reveal any significant environment dependence in the  $f_Q$  evolution. Therefore, the potential contamination in our environment measurements cannot explain the lack of environment dependence in the  $f_Q$  evolution.

Many physical mechanisms have been proposed to explain the quenching of star-formation observed in our universe. Often star-formation quenching is classified into two distinct mechanisms: mass-dependent and environment dependent mechanisms (Baldry et al. (2006); ?). From the mass-dependence presented in all  $f_Q$  in Figure 4, we concur that galaxy stellar mass plays a significant role in quenching. Furthermore, since  $f_Q$  in high density environments is always greater than  $f_Q$  in low density environments, our results also reflect the environment dependence in quenching.

## 5. DISCUSSION

## 6. SUMMARY

Using a stellar mass complete galaxy sampled derived from SDSS and PRIMUS accompanied by a consistently measured galaxy environment from robust spectroscopic redshifts, we measure the stellar mass functions for star-forming and quiescent in low and high density environments over the redshift range  $z < 0.8$ . From these stellar mass functions, we compare the proportion of quiescent galaxies within the subsamples by computing quiescent fraction for each of them. Afterwards, in order to better quantify the evolution of the quiescent fraction over cosmic time, we fit our quiescent fraction anchored at a fiducial mass. From our analyses we find the following notable results.

1. By noting the SMF evolution for each of our subsamples, we loosely trace the evolutionary tracks our galaxy population tracing the general evolutionary tracks of quenching mechanisms and galaxies infalling into high density environments.
2. From the SMFs, we find that our galaxy population in high density environments, both star-forming and quiescent, have a higher median mass, thus confirming the mass density relation and mass-segregation in different environments.
3. In all of our subsamples, we find a clear mass dependence in  $f_Q$ .  $f_Q$  increases monotonically with galaxy stellar mass.
4. Both from the  $f_Q$  evolution and  $f_Q(\mathcal{M}_{\text{fid}})$  values,  $f_Q$  increases with redshift.
5. We conclude, by fitting our  $f_Q$ , that  $f_Q$  in high density environments is greater than  $f_Q$  in low density environments regardless of mass and redshift. This result confirms the environmental dependence of quenching mechanisms stated previously in literature.

TABLE 2  
ENVIRONMENT DEFINING APERTURE DIMENSIONS

Radius (Mpc)	Height (Mpc)	Edge Effect Cutoff	High Env (galaxies)	Low Env (galaxies)
1.0	50	80%	1.5	0.0
2.0	50	75%	4.0	0.0

6. Comparison of the  $f_Q(\mathcal{M}_{\text{fid}})$  evolution for high and low density environments that the total evolution is independent of environment.

We thank people...

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#### APPENDIX