

1 Report

by Gullik Vetvik Killie, Steffen Brask Add yourself

Abstract

Lorem ipsum dolor sit amet, consectetur adipisicing elit, sed do eiusmod tempor incididunt ut labore et dolore magna aliqua. Ut enim ad minim veniam, quis nostrud exercitation ullamco laboris nisi ut aliquip ex ea commodo consequat. Duis aute irure dolor in reprehenderit in voluptate velit esse cillum dolore eu fugiat nulla pariatur. Excepteur sint occaecat cupidatat non proident, sunt in culpa qui officia deserunt mollit anim id est laborum.

1.1 Introduction

Langmuir probes, as well as other instruments, are often situated outside of space crafts. The presence of the spacecraft causes disturbances to the local conditions we want to measure, these disturbances need to be corrected for. Many studies have been done on the flow of plasma around spacecraft and of the wake the spacecraft creates (miloch'wake'2010; engwall'wake'2006).

1.2 Theory

1.3 Numerical Methods

1.3.1 Numerical methods

To solve the problem numerically we use the EMSES code. EMSES uses the standard PIC method for plasma simulations. In the code we are able to define a spacecraft body, and the code then calculates the potential on that body using the capacitance matrix method. Although EMSES has the capability to do a full electromagnetic calculation, we have opted to use the poisson's equation solver for electrostatic problems. In the EMSES system we can define sunlit surfaces based upon an angle, and a current density. Sunlit surfaces will then emit electrons based upon a energy distribution. For a complete description of EMSES' capabilities see (nakashima'ohhelp:'2009) Parameters are chosen to simulate the sun at the earth, but with an enhanced flux to emphasize the effect in question.

1.3.2 Theoretical calculations

$$\begin{aligned}
I_i &= \begin{cases} A|q|n_\infty \sqrt{\frac{8kT_i}{\pi m_i}} (\text{without flow}) \\ \frac{1}{6}A|q|n_\infty V + \frac{5}{6}A|q|n_\infty \sqrt{\frac{8kT_i}{\pi M_i}} (\text{with flow}) \end{cases} \\
I_e &= -A|q|n_\infty \sqrt{\frac{8kT_e}{\pi m_e}} \exp\left(\frac{|q|\Phi_d}{kT_e}\right) \\
I_{ph} &= \frac{1}{6}AJ_s
\end{aligned} \tag{1.1}$$

Submitting these equations to $I_i + I_e + I_{ph} = 0$ and solving for Φ_d , we can get

$$\begin{aligned}
\Phi_d &= \frac{kT_e}{|q|} \ln\left(\sqrt{\frac{T_i}{T_e}} \sqrt{\frac{m_e}{m_i}} + \frac{J_s}{6n_\infty|q|} \sqrt{\frac{\pi m_e}{8kT_e}}\right) + \Phi_0 \\
\Phi_d &= \frac{kT_e}{|q|} \ln\left(\frac{5}{6} \sqrt{\frac{T_i}{T_e}} \sqrt{\frac{m_e}{m_i}} + \frac{V}{6} \sqrt{\frac{\pi m_e}{8kT_e}} + \frac{J_s}{6n_\infty|q|} \sqrt{\frac{\pi m_e}{8kT_e}}\right) \\
\Phi_0 &= kT_{ph}, \quad V = \text{speed of flow}
\end{aligned} \tag{1.2}$$

The theory is based on thin-sheath approximation, so if SC size is much larger than Debye length of the photoelectron, we can use it.

	Without flow	With flow
Without Photon emission	-0.893	-0.792
With Photon emission	-0.718	-0.662

From these calculations we can expect the potential over the spacecraft to rise in the cases where we have photon emission.

1.3.3 Test case setup

We wish to simulate the effects of Photon emitted electrons in different test cases, and have thus set up the following 6 cases:

Case	Plasma flow	Photon emission
1:	41600 \vec{e}_x m/s	0
2:	-41600 \vec{e}_z m/s	0
3:	-41600 \vec{e}_y m/s	0
4:	41600 \vec{e}_x m/s	$-10^{-3} A/m^3 \vec{e}_x$
5:	-41600 \vec{e}_z m/s	$-10^{-3} A/m^3 \vec{e}_x$
6:	-41600 \vec{e}_y m/s	$-10^{-3} A/m^3 \vec{e}_x$

So test case 1-4, 2-5, and 3-6 are the “same” cases except that we run the simulation with and without photon emission to compare the cases two and two. We define 3 geometric objects, the spacecraft itself and two probes. In all cases the B field is in the \vec{e}_z direction.

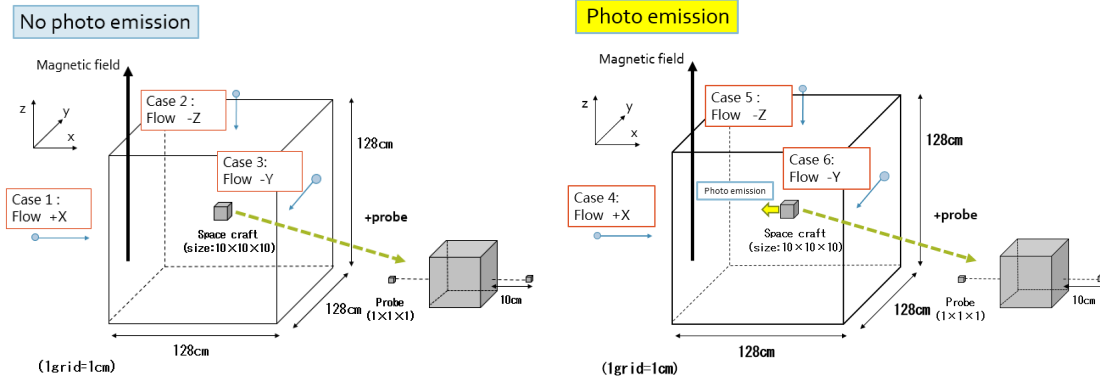


Figure 1.1: Research about spacecraft and surroundings without P-E and without. Left figure is no P-E simulation cases. Right figure is P-E simulation cases.

1.4 Results

1.4.1 Induced electric current

The plasma is flowing in relation to the coordinate system in the simulations. Due to this an induced electrical field, $\vec{\varepsilon}$, will appear. The induced electrical field will neutralize the Lorentz force. Combined with the electrostatic approximation we can obtain the $\vec{\varepsilon}$

$$\vec{\varepsilon} = \vec{v}_D \times \vec{B} \quad (1.3)$$

This will cause a potential gradient perpendicular to the plasma flow and the magnetic field, using the electrostatic approximation we obtain the magnitude of the gradient.

$$\int E dx = -\phi \quad (1.4)$$

$$\phi = - \int \vec{v}_d \times \vec{B} \approx - \int (41600 \text{ m/s} \cdot 50 \text{ E} - 6 \text{ T}) dx \quad (1.5)$$

$$|\nabla \phi| = 2.08 \text{ m}^{-1} \quad (1.6)$$

Figure 1.2 shows the measured potential at case 6.

1.4.2 Photoemission paths

The electrons emitted from the spacecraft due to the photoelectric effect, have a kinetic energy corresponding to a Maxwellian distribution with a temperature of $T_{ph} = 3.8481 \cdot 10^4 \text{ K}$. Figure 1.3 illustrates the trajectories of the emitted electrons in simulation 6. As the probes are situated 10cm to the sides of the spacecraft on the x -axis, the probes may be hit by the photo-emitted electrons. In the following section, ??, we show the number of electrons hitting the probes.

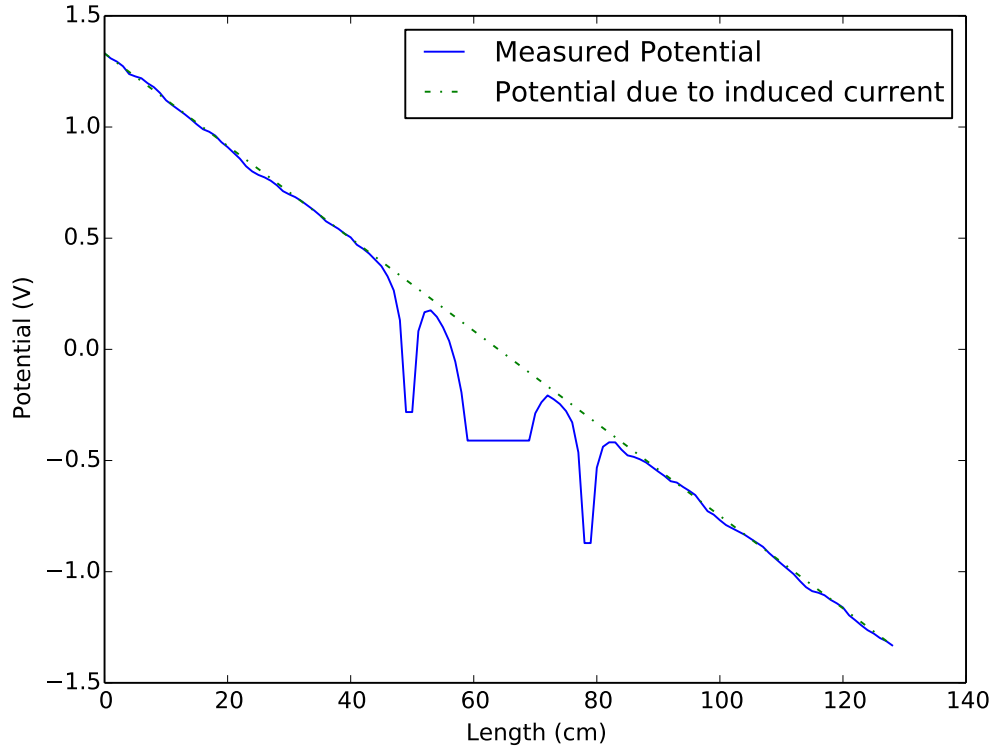


Figure 1.2: The blue line is the potential along direction x for simulation 6. In this case the potential gradient is along the x -axis. The dotted green line is the potential caused by the induced electrical field. This should be accounted for if we want to find the potential at the spacecraft and the probes.

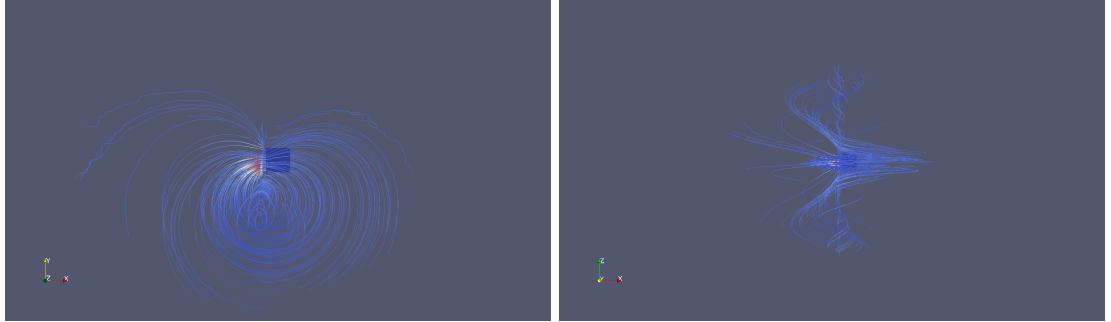


Figure 1.3: The trajectories of the electrons emitted by the photoelectric effect in simulation 6. The possible paths of the photoemitted electrons coincide with the volume occupied by the langmuir probes. The photoemitted electrons are strongly affected by the magnetic field \vec{B} , and follows a gyrating path guided by \vec{B} . The photoemitted electrons are in all the studied cases emitted from the spacecraft in $-x$ direction, and the paths are similar. The langmuir probes are situated 10cm to each side of the spacecraft along the x -direction. (NOTE, should have axis labels, and domain length.)

1.4.3 Accumulated Photoelectrons on Langmuir Probes

In the simulations both electrons and photoelectrons are absorbed by the probes. Table 1.4 shows the current of both regular electrons, as well as photoelectrons, interacting with the probes. The photoelectrons are emitted on the left side of the spacecraft and we see a larger current of photoelectrons here. The current caused by the photoemission is 10 – 100 smaller than the current from the electrons in the plasma.

	probe to +x	probe to -x
case 1	2.207E-07	2.851E-07
case 2	2.794E-07	2.807E-07
case 3	2.786E-07	2.772E-07
case 4	2.117E-07	2.609E-07
case 5	2.754E-07	2.616E-07
case 6	2.777E-07	2.512E-07

	probe to +x	probe to -x
case 4	2.3324E-09	3.49699E-08
case 5	1.9598E-09	3.51036E-08
case 6	1.3824E-09	3.66922E-08

Figure 1.4: The left table shows the electron current hitting the inserted probes at for the simulated cases 1 – 6. On the table to the right the photoelectrons hitting the probes are shown. The photoelectron current is smaller than the electron current from the plasma, varying from 10 – 100 times smaller.

1.4.4 Potential difference with P-E and no P-E

Case 1 vs case 4

In case 4 we have the emitted electrons in the negative flow direction. As expected this leads to a drop in potential in the left probe which is facing the plasma flow. The potential drop over the probe is 0.02V or 3.8%. The right probe is now the wake where we have a drop in the ion density. This yields a large drop in potential compared to the left probe, but it also has a larger potential drop than the left probe when comparing case 1 and 4. This might be because the increase in electron density on the left side redirects more ions from the right side. The potential drop in the right probe comparing the two cases is 0.07V or 10%. The potential rise over the satellite is 0.17V or 28%.

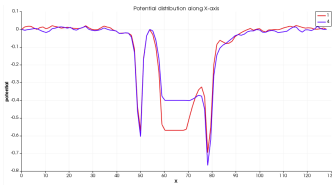


Figure 1.5: Potential of satellite and surroundings in the x-direction for case 1 and case 4. Note that the potential drop is much greater in the right probe when comparing the two cases.

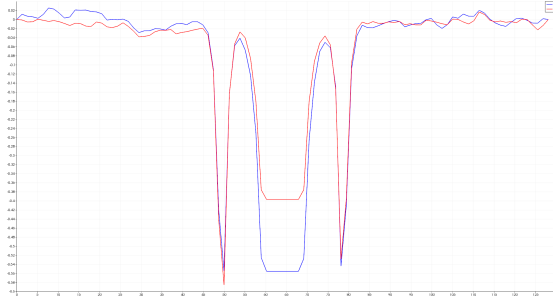


Figure 1.6: Figure show potential of satellite and surroundings in the x-direction for case 2 and 5. Note that we have a small potential drop in the left probe but a potential rise in the right probe when comparing the two cases.

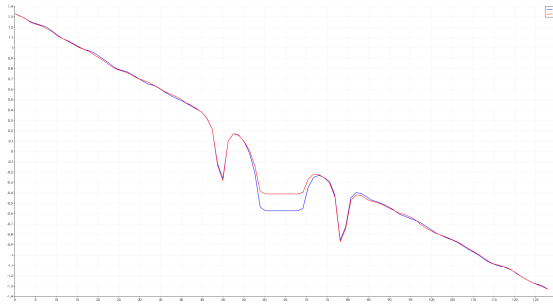


Figure 1.7: Figure show potential of satellite and surroundings in the x-direction for case 3 and 6. Note that there is almost the same drop in potential for both probes.

Case 2 vs case 5

With the flow of emitted electrons in the negative x direction we would expect a rise in electron density around the left probe. This can be seen in figure 1.6 where we see a potential drop of 0.03V or 5.4% over this probe compared to case 5 with no emitted electrons. On the right probe we have a small rise in the potential of 0.018V or 3.3%. With no emitted electrons on this side of the satellite the rise in potential can be explained by looking at the increase in ion density as seen in figure (need ref). In the satellite we have a potential rise of 0.16V or 28%. So the potential change in the probes are small compared to the potential change in the satellite.

Case 3 vs case 6

Here we have almost the same drop in potential over both probes of about 0.03V. However this change yields 12% for the left probe but only 3% for the right probe. Over the satellite we have a potential rise of 0.15V which yields a 26% change.

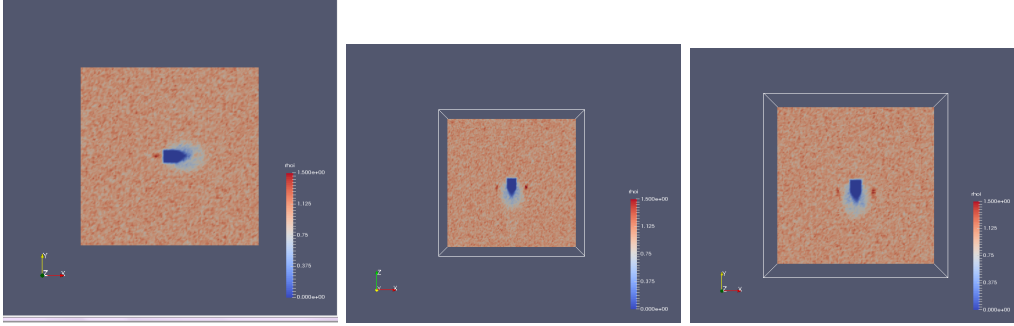


Figure 1.8: Ion density of spacecraft and surroundings without P-E Figure on the left displays case 1. Middle figure displays case 2. Rightmost figure displays case 3.

1.4.5 Wake plots

The spacecraft is moving relative to the plasma flow, this causes a wake to be formed in the vicinity of the craft. Figures 1.8 illustrates the flow in case 1, 2 and 3.

1.5 Discussion

1.6 Conclusions

In the simulated cases with exaggerated

- Proposal for further studies (Probably see if photoemmission is relevant in tenuous plasma (MEO CASE, magnetospheric tail lobes))