

Cosmological Perturbations

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Most of the conventions and notations in [1, ch. 5] will be followed.
Suppose the metric can be expanded up to the linear order as

$$g = g^{(0)} + \epsilon g^{(1)} + O(\epsilon^2). \quad (1)$$

The background metric $g^{(0)}$ takes the Robertson–Walker form

$$g_{\mu\nu}^{(0)} dx^\mu dx^\nu = -N^2(t) dt^2 + a^2(t) d\Omega_{3F}^2, \quad (2)$$

in which $d\Omega_{3F}^2 = d\chi^2 + \chi^2(d\theta^2 + \sin^2\theta d\phi^2)$ is the dimensionless flat spatial metric. The linear perturbation can be decomposed into scalar, vector and tensor parts

$$g_{00}^{(1)} = -N^2 E, \quad (3)$$

$$g_{i0}^{(1)} = g_{0i}^{(1)} = Na(F_{,i} + G_i), \quad (4)$$

$$g_{ij}^{(1)} = a^2(A\delta_{ij} + B_{,i,j} + C_{i,j} + C_{j,i} + D_{ij}). \quad (5)$$

Here one has some weird condition, where the contractions do not follow the one-up-one-down tradition

$$C_{i,i} = G_{i,i} = 0, \quad D_{ij,i} = 0, \quad D_{ii} = 0. \quad (6)$$

1 Transform under diffeomorphism

Consider a diffeomorphism generated by $\xi^\mu = O(\epsilon)$, which can in turn be decomposed into $\xi_0 = \zeta$, $\xi_i = \xi_{i,i}^S + \xi_i^V$.

One has here again some weird condition, where the contractions do not follow the one-up-one-down tradition

$$\xi_{i,i}^V = 0. \quad (7)$$

The Lie derivative of the metric $\mathbb{L}_\xi g$ is

$$(\mathbb{L}_\xi g)_{\mu\nu} = \xi^\lambda g_{\mu\nu,\lambda} + \xi^\lambda_{,\mu} g_{\lambda\nu} + \xi^\lambda_{,\nu} g_{\mu\lambda}. \quad (8)$$

In components and expansion, these are

$$(\mathbb{L}_\xi g)_{00} = 2\dot{\zeta} - 2\zeta \frac{\dot{N}}{N} + O(\epsilon^2), \quad (9)$$

$$(\mathbb{L}_\xi g)_{i0} = (\mathbb{L}_\xi g)_{0i} =, \quad (10)$$

$$(11)$$

2 Scalar perturbations

3 Vector perturbations

4 Tensor perturbations

References

- [1] Steven Weinberg. *Cosmology*. Oxford University Press, 2008. ISBN: 9780198526827.
URL: <https://global.oup.com/academic/product/cosmology-9780198526827>.