



Lecture 14: The Sun

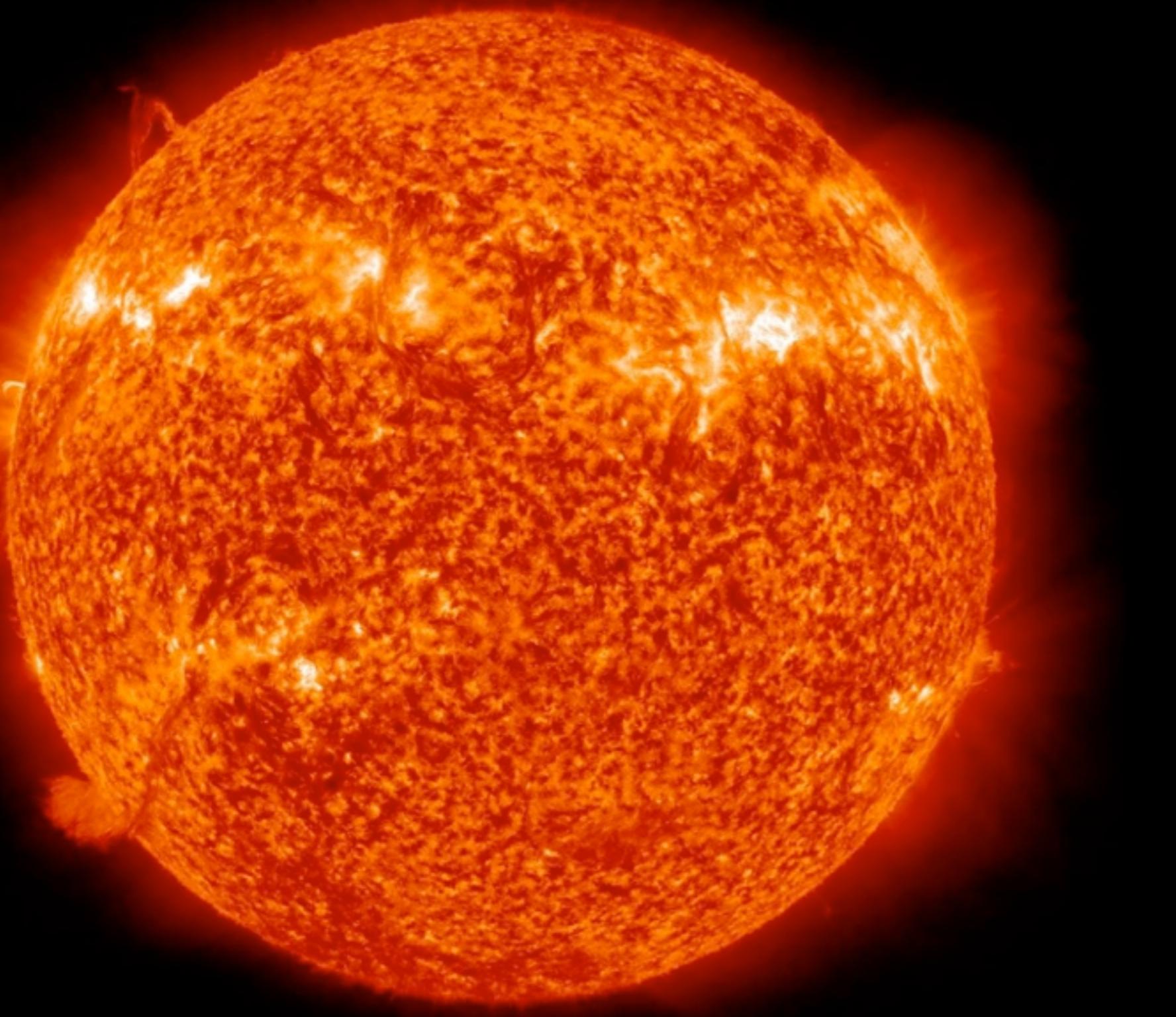


Astronomy

with your host:



Coop



$m \sim 2\text{e}30 \text{ kg}$
 $\sim 300,000 m_e$

$r \sim 7\text{e}5 \text{ km}$
 $\sim 110 r_e$

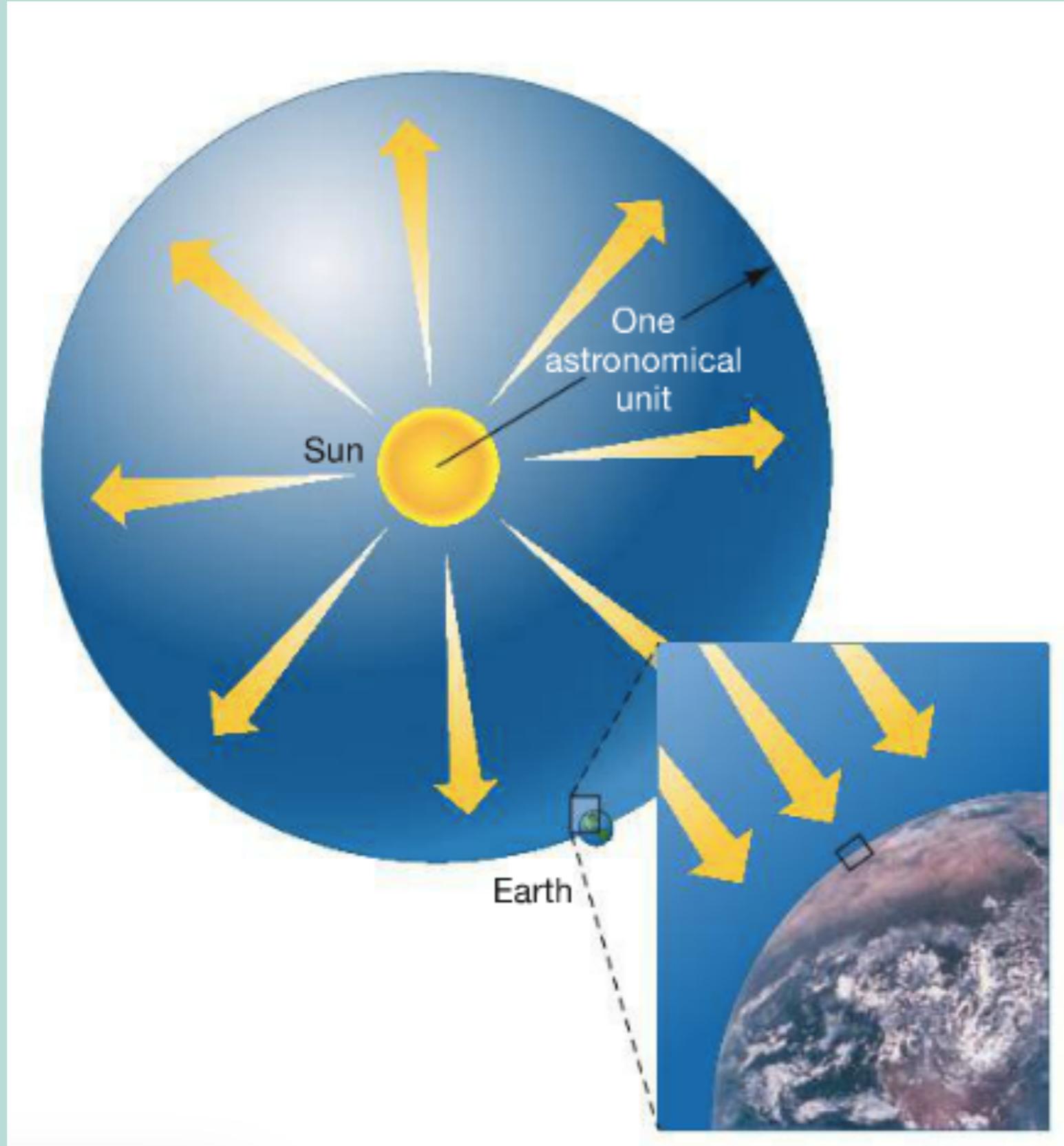
$d \sim 28,000$
years

density ~
 1400 kg/m^3

The Sun: Ch 16

- 0.5 degrees in the sky
- Differential Rotation
- 25 day rotation period at the equator, 36 at the poles
- Effective temperature of 5800 K

The Sun: Ch 16



Just how bright is the sun?

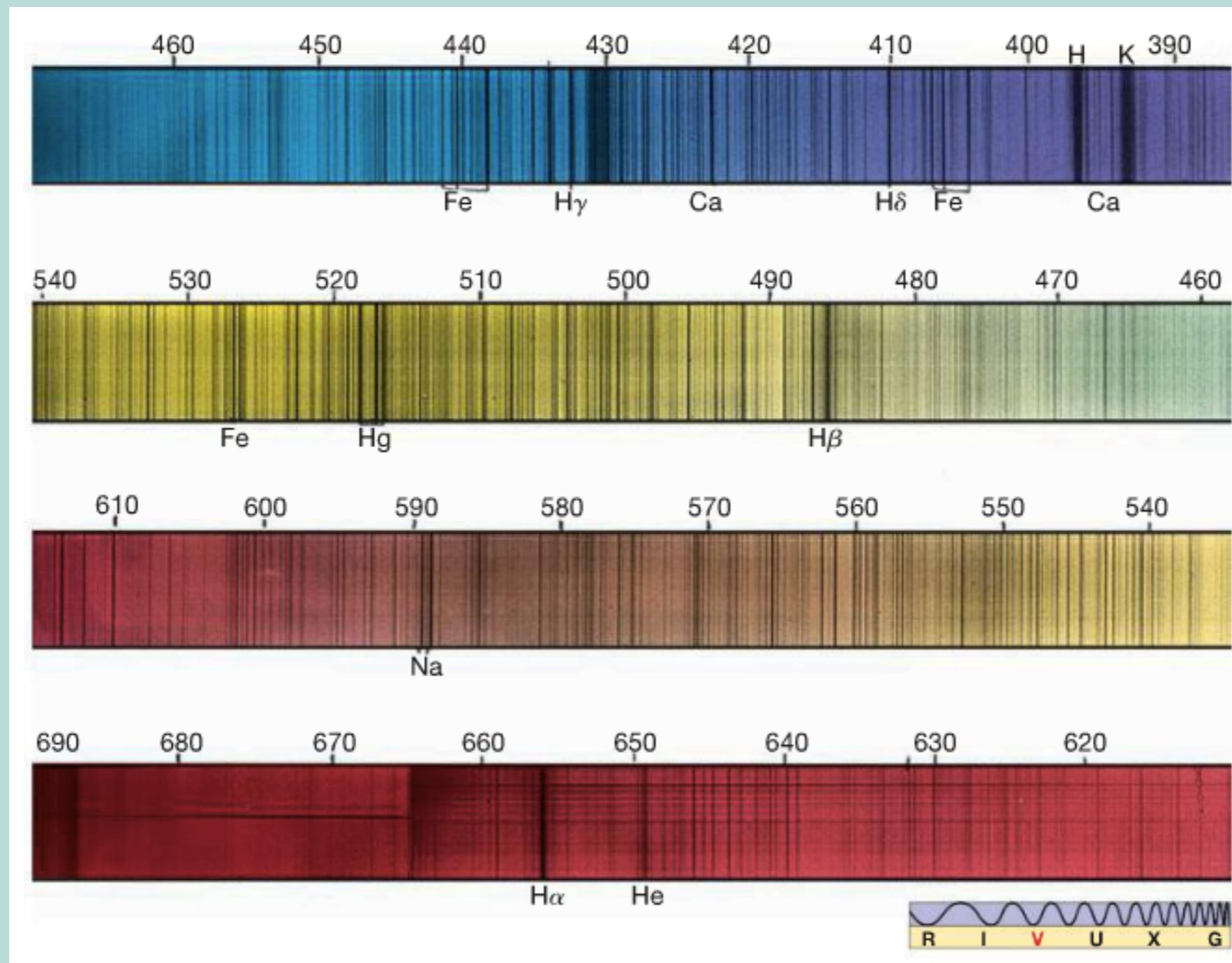
- Solar constant: 1400 W/m^2
- 50% reaches earth, 30% goes to the atmosphere, 20% reflected away
 - About a mol of square meters in an AU sphere around the sun
 - Result: $4 \times 10^{26} \text{ W}$
 - 10 billion 1 megaton bombs/second
 - Could evaporate the world's oceans in 6 seconds, melt the crust in 3 minutes

The Sun: Ch 16

TABLE 16.2 The Composition of the Sun

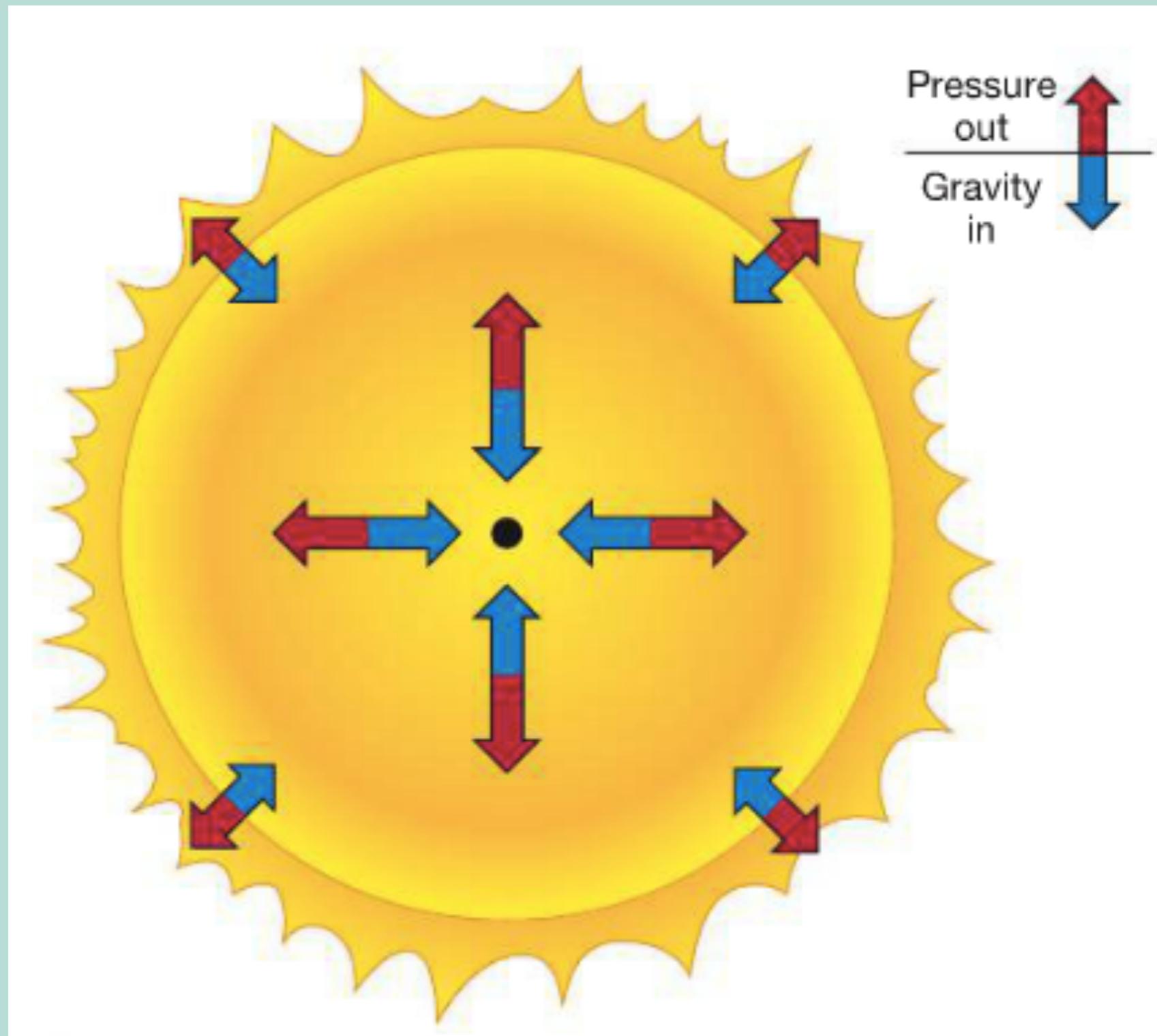
Element	Percentage of Total Number of Atoms	Percentage of Total Mass
Hydrogen	91.2	71.0
Helium	8.7	27.1
Oxygen	0.078	0.97
Carbon	0.043	0.40
Nitrogen	0.0088	0.096
Silicon	0.0045	0.099
Magnesium	0.0038	0.076
Neon	0.0035	0.058
Iron	0.0030	0.14
Sulfur	0.0015	0.040

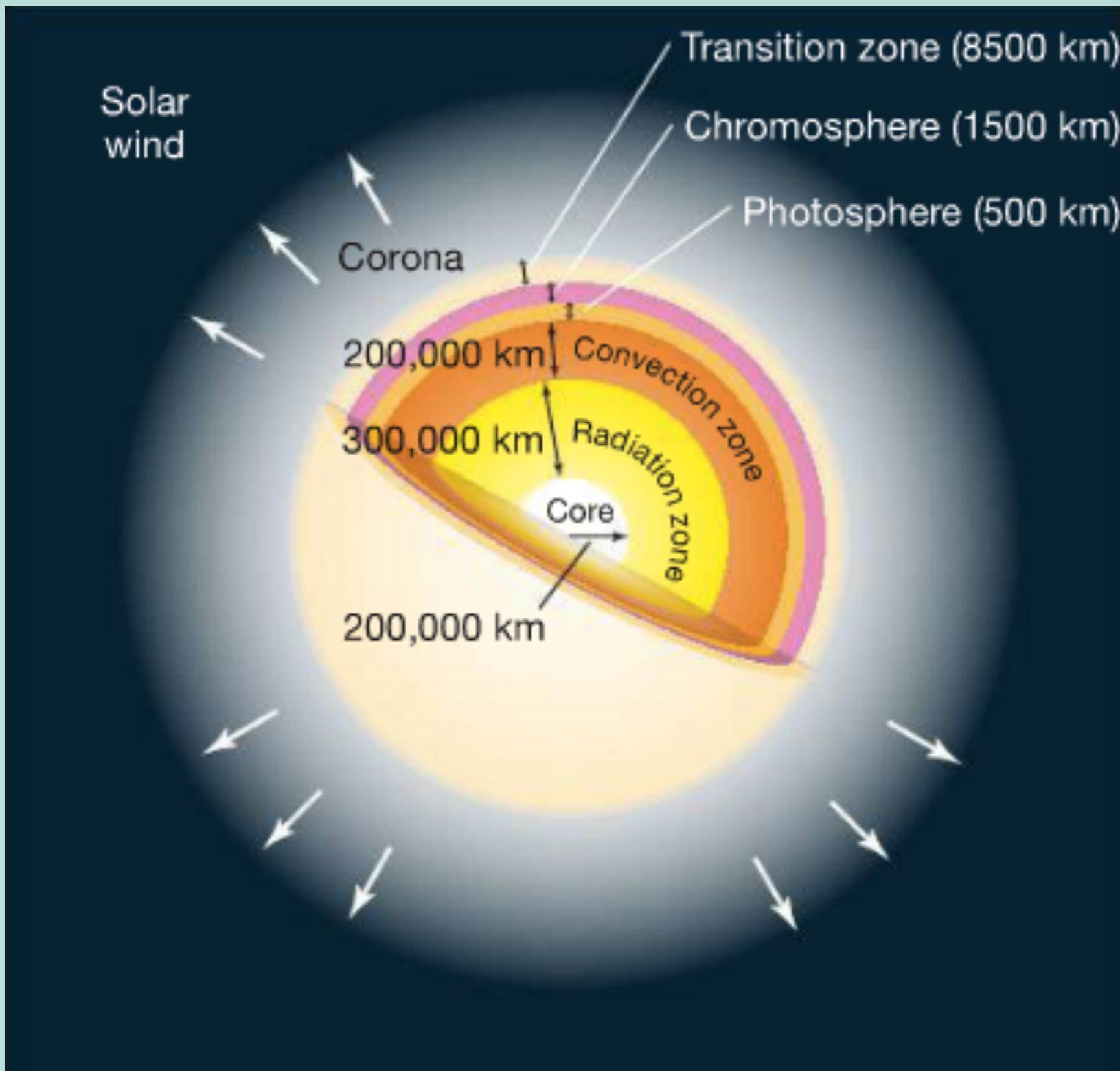
Solar Spectrum



The Sun: Ch 16

- Recall, we learn not just about the emitted radiation (emission lines), but also the intermediate atmosphere of the sun (absorption lines)





The Sun: Ch 16

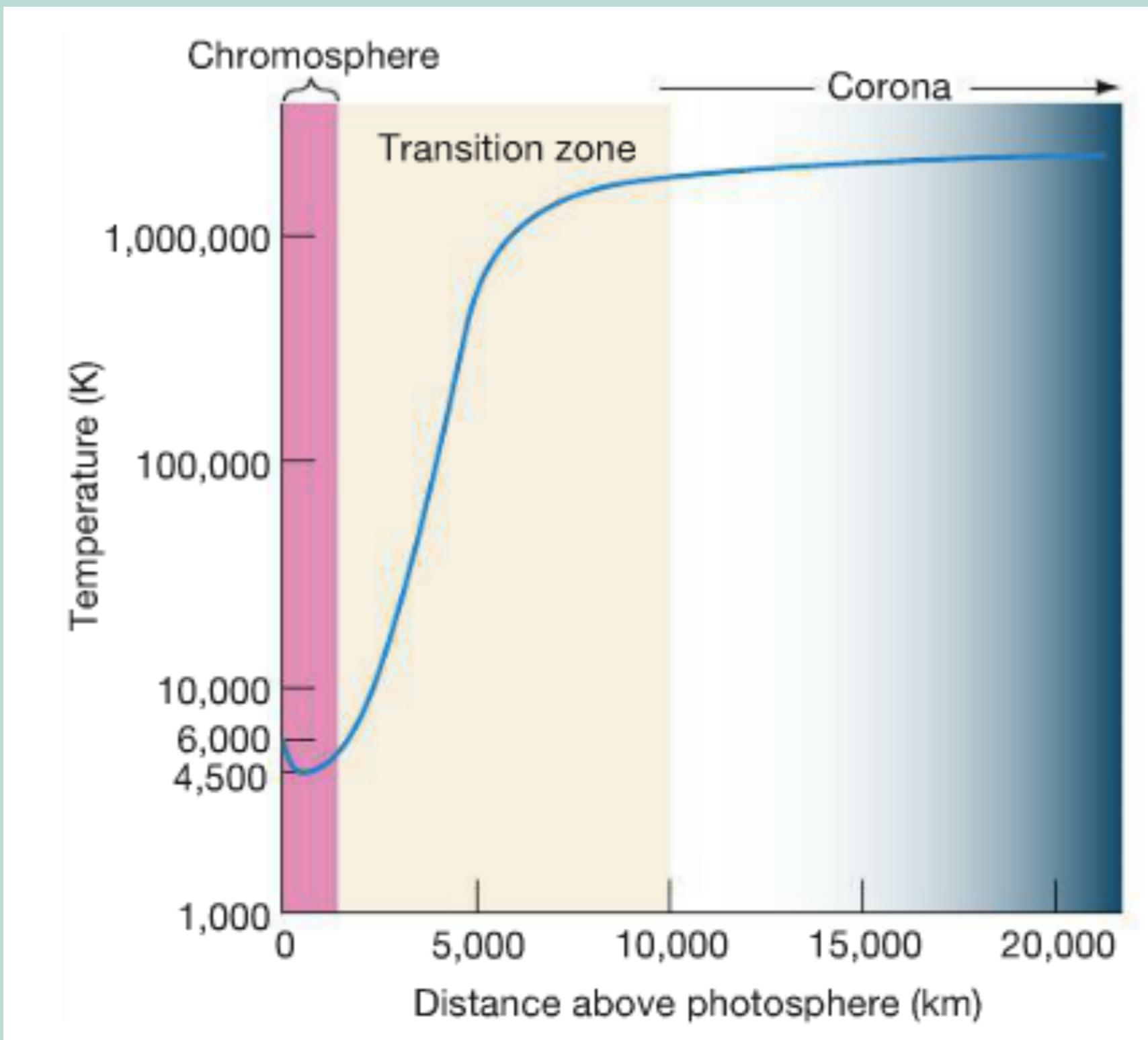
- In the core it's too hot to form atoms, so radiation doesn't get absorbed by anything (radiation zone)
- Further out, it's cool enough to form atoms (how cool?) and so ALL the light gets absorbed and turned into heat and reflected light
- This heat CONVECTS to the surface which glows like a black body (photosphere)

The Sun: Ch 16

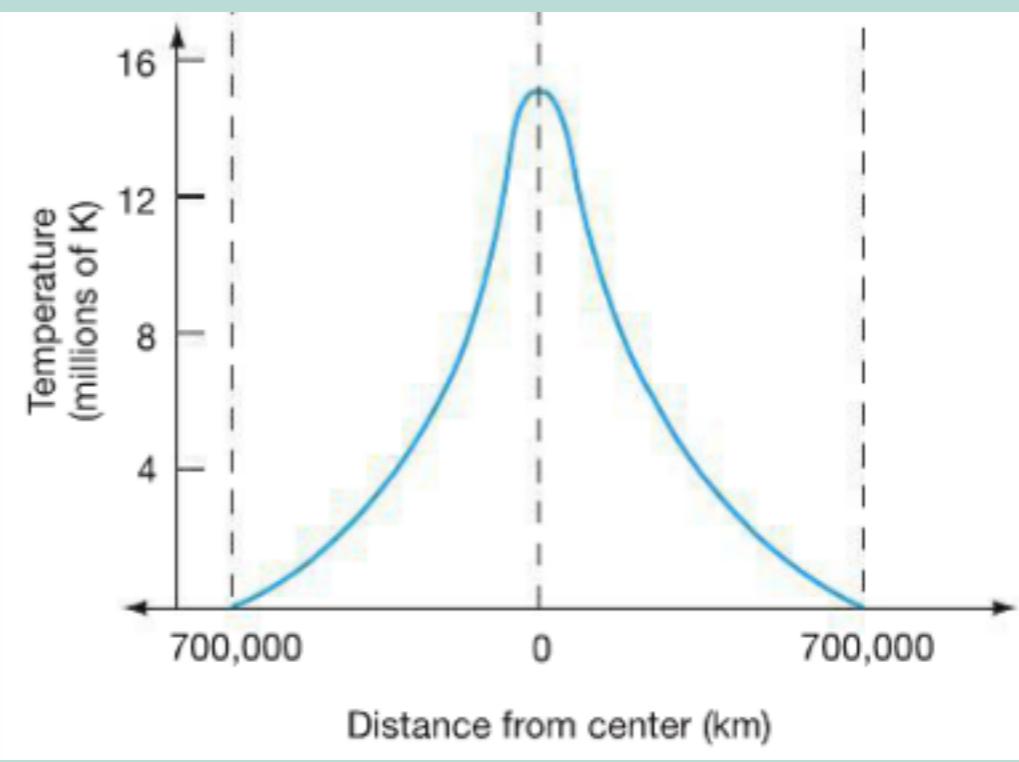
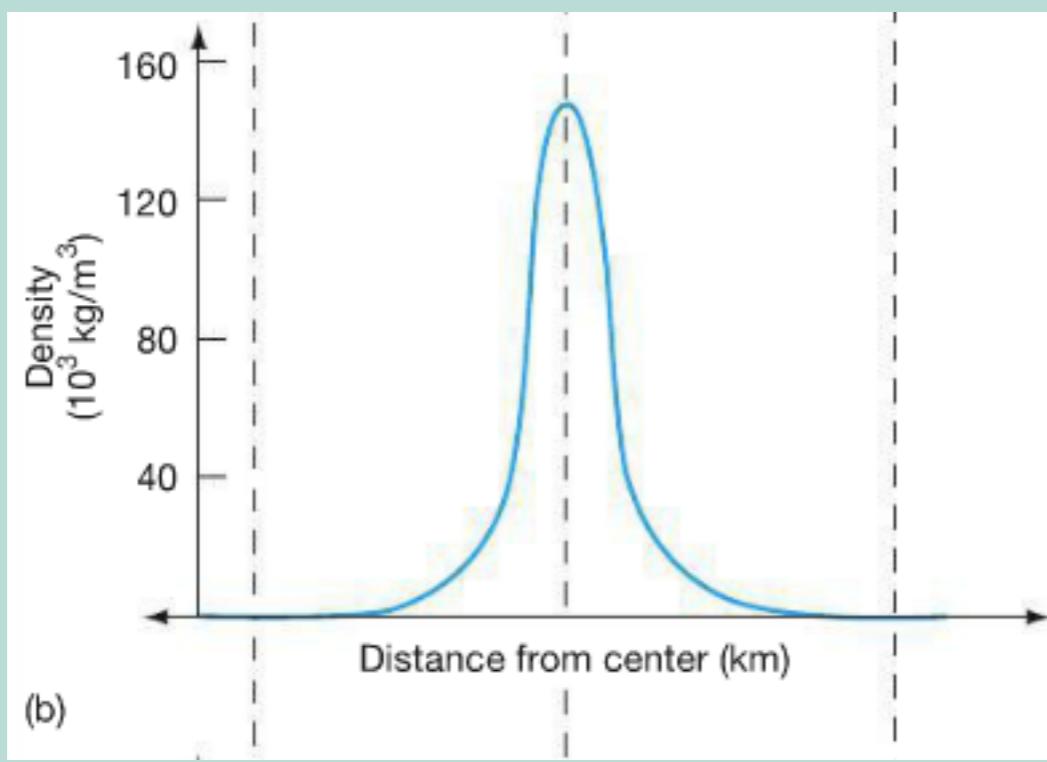
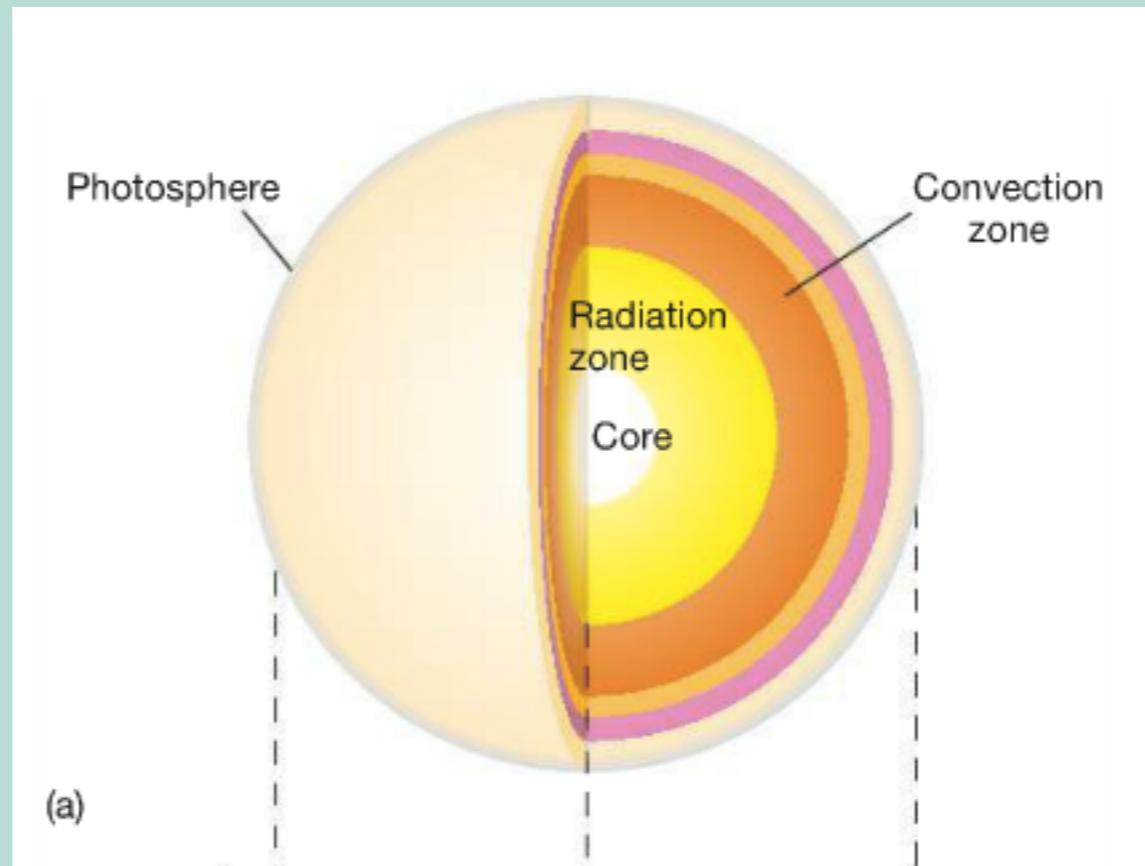
TABLE 16.1 The Standard Solar Model

Region	Inner Radius (km)	Temperature (K)	Density (kg/m ³)	Defining Properties
Core	0	15,000,000	150,000	Energy generated by nuclear fusion
Radiation zone	200,000	7,000,000	15,000	Energy transported by electromagnetic radiation
Convection zone	496,000*	2,000,000	150	Energy carried by convection
Photosphere	696,000*	5800	2×10^{-4}	Electromagnetic radiation can escape—the part of the Sun we see
Chromosphere	696,500*	4500	5×10^{-6}	Cool lower atmosphere
Transition zone	698,000*	8000	2×10^{-10}	Rapid increase in temperature
Corona	706,000*	3,000,000	10^{-12}	Hot, low-density upper atmosphere
Solar wind	10,000,000	> 1,000,000	10^{-23}	Solar material escapes into space and flows outward through the solar system

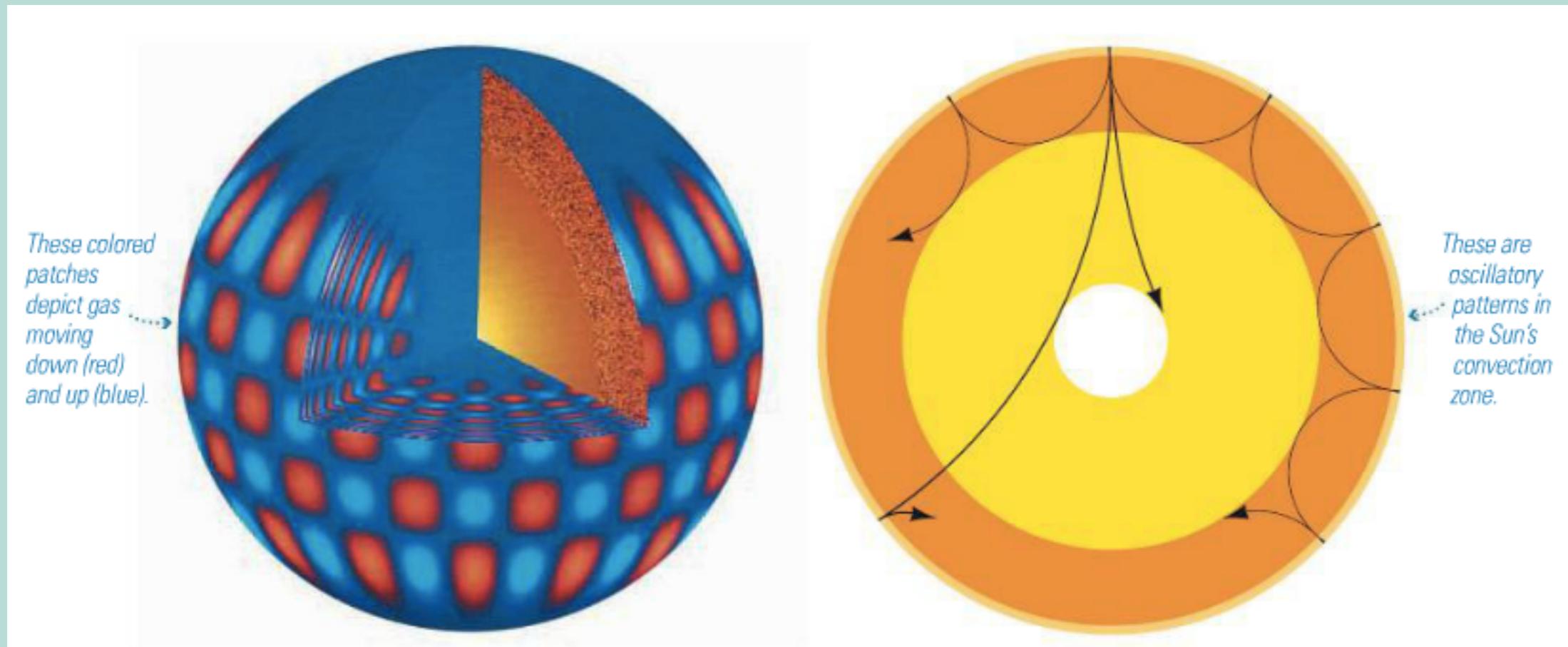
Atmospheric Physics is Hard



The Sun: Ch 16

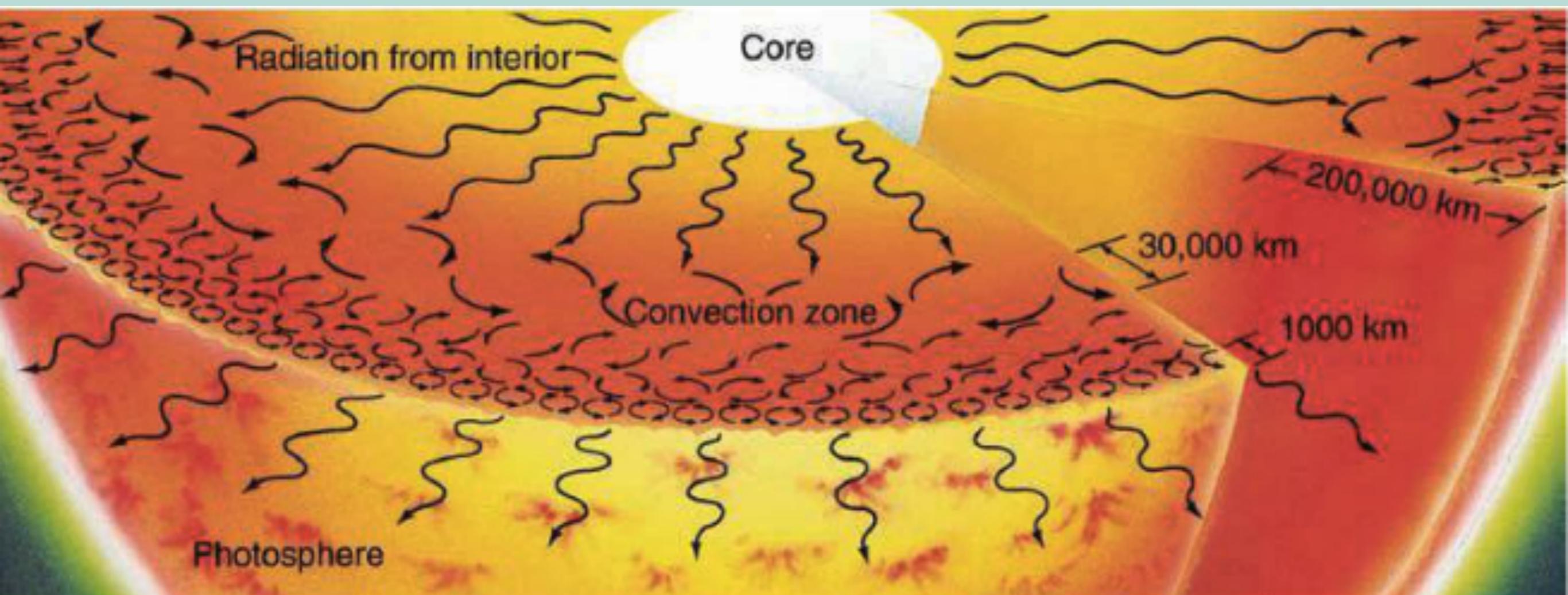


Helioseismology



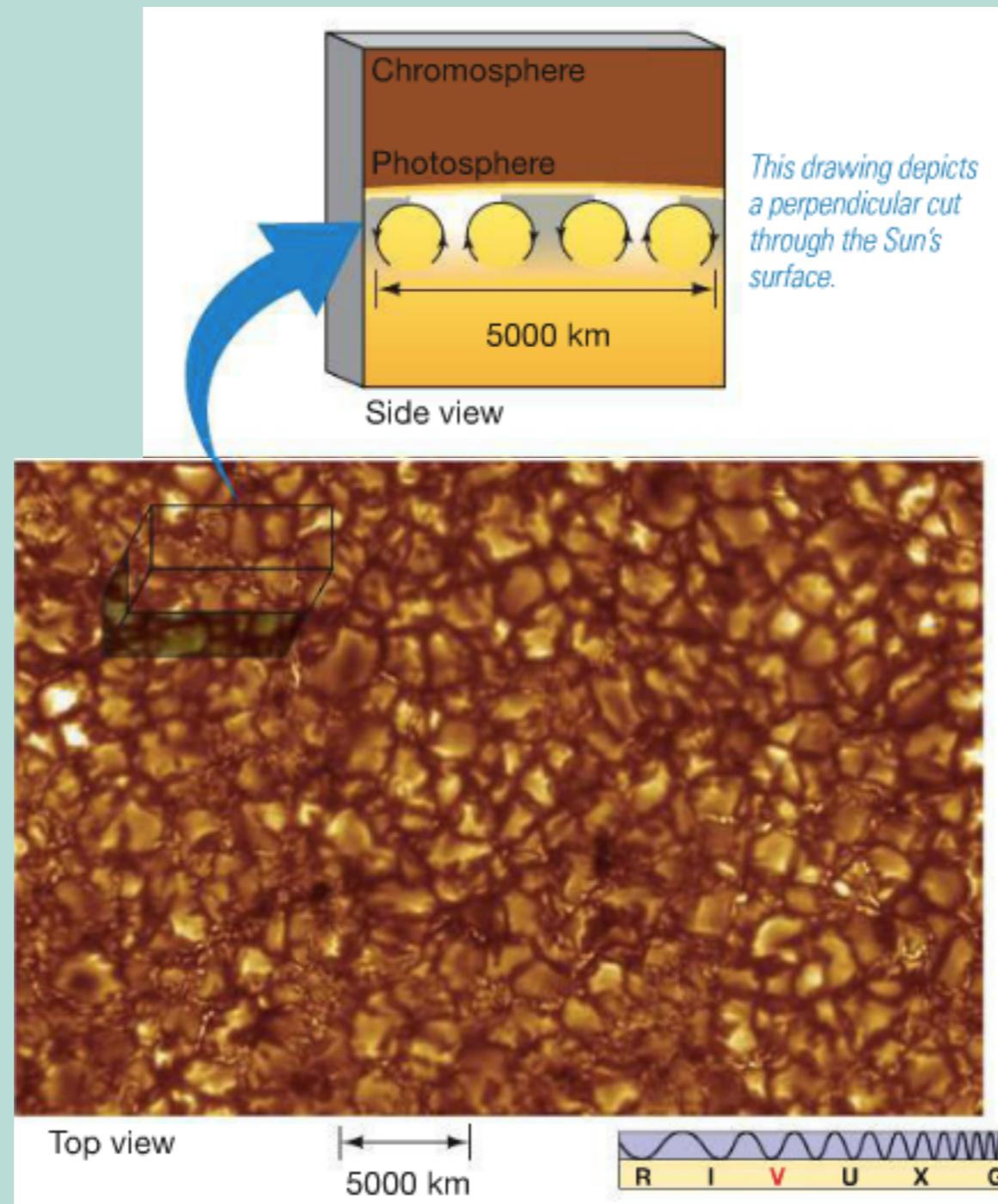
Global Oscillations Network Group (GONG)

Convection Cells



Hierarchy of cells

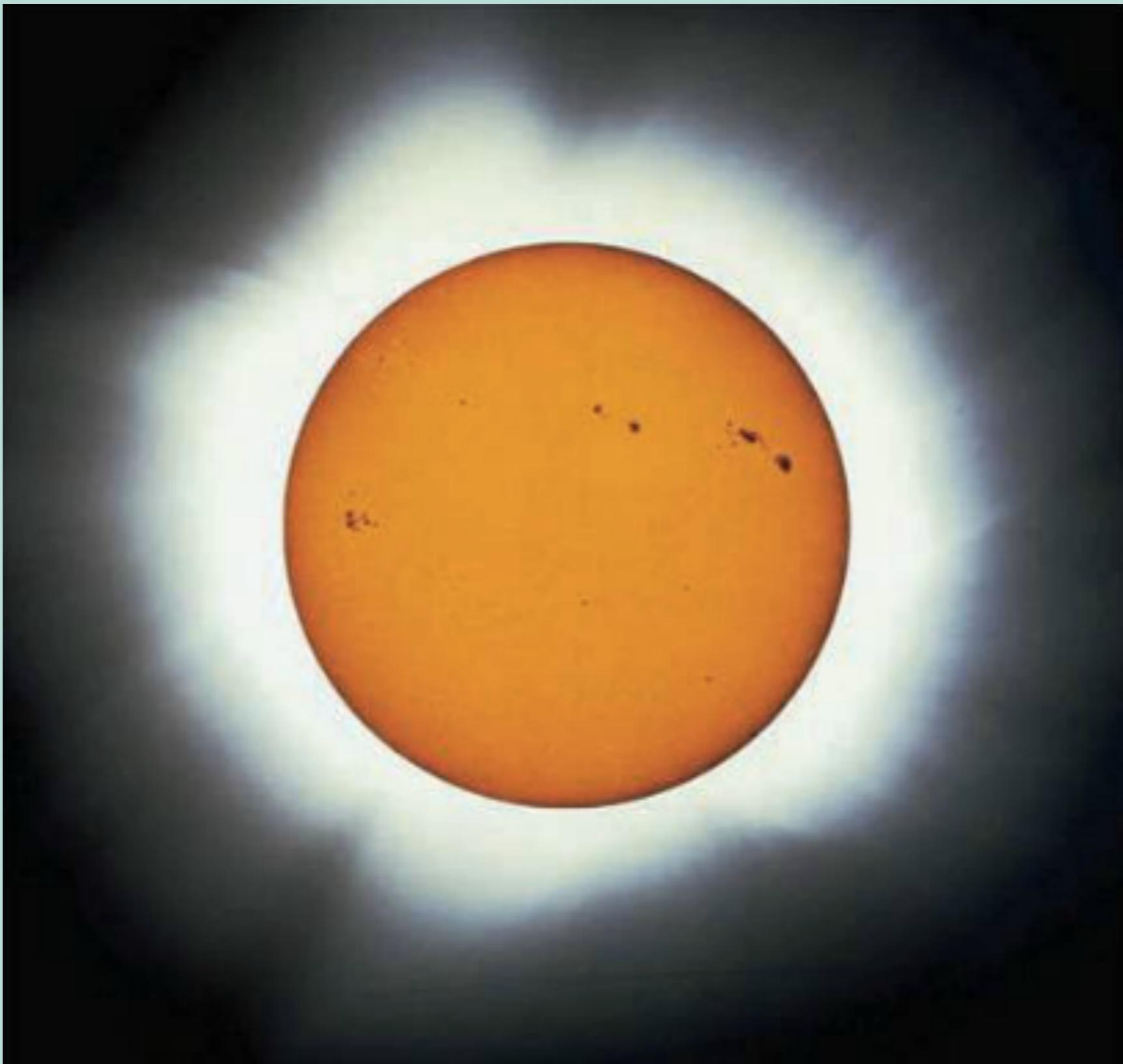
Solar Granules



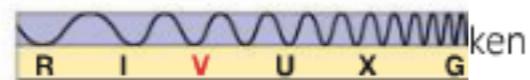
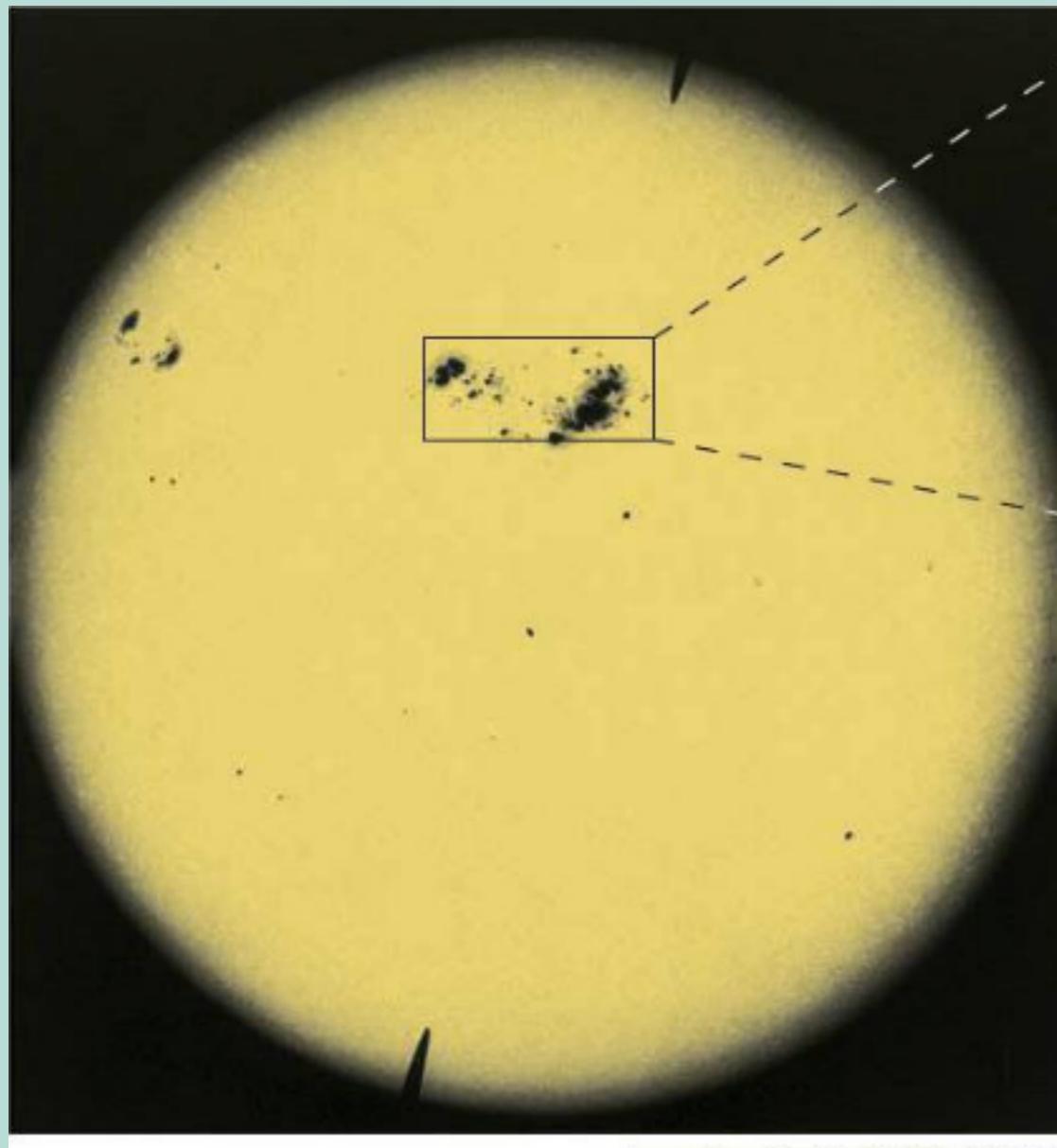
- Each granule is approximately the size of the Earth!

The Sun: Ch 16

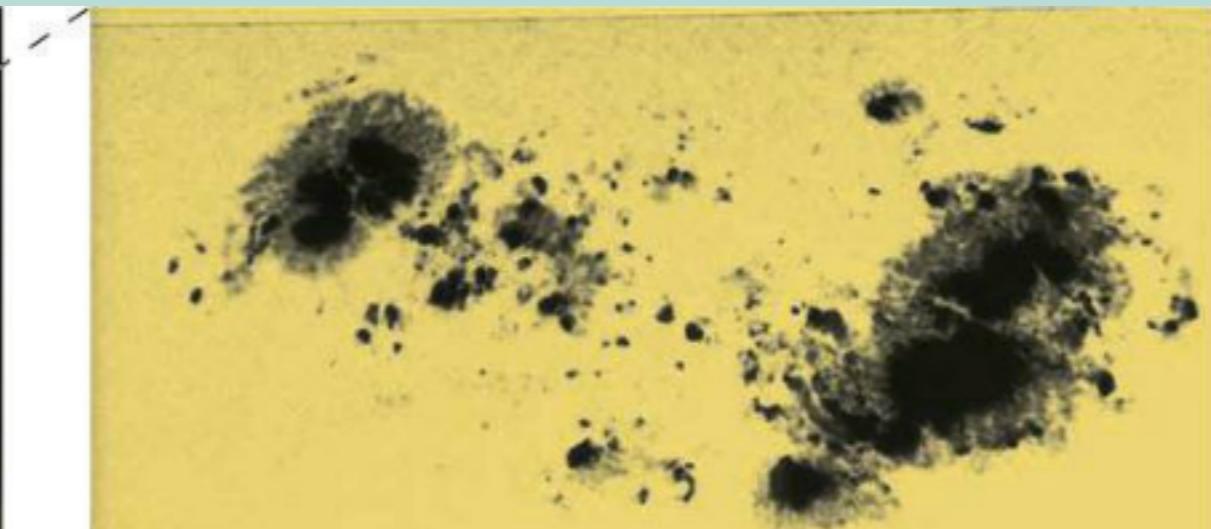
Sun Spots from Earth



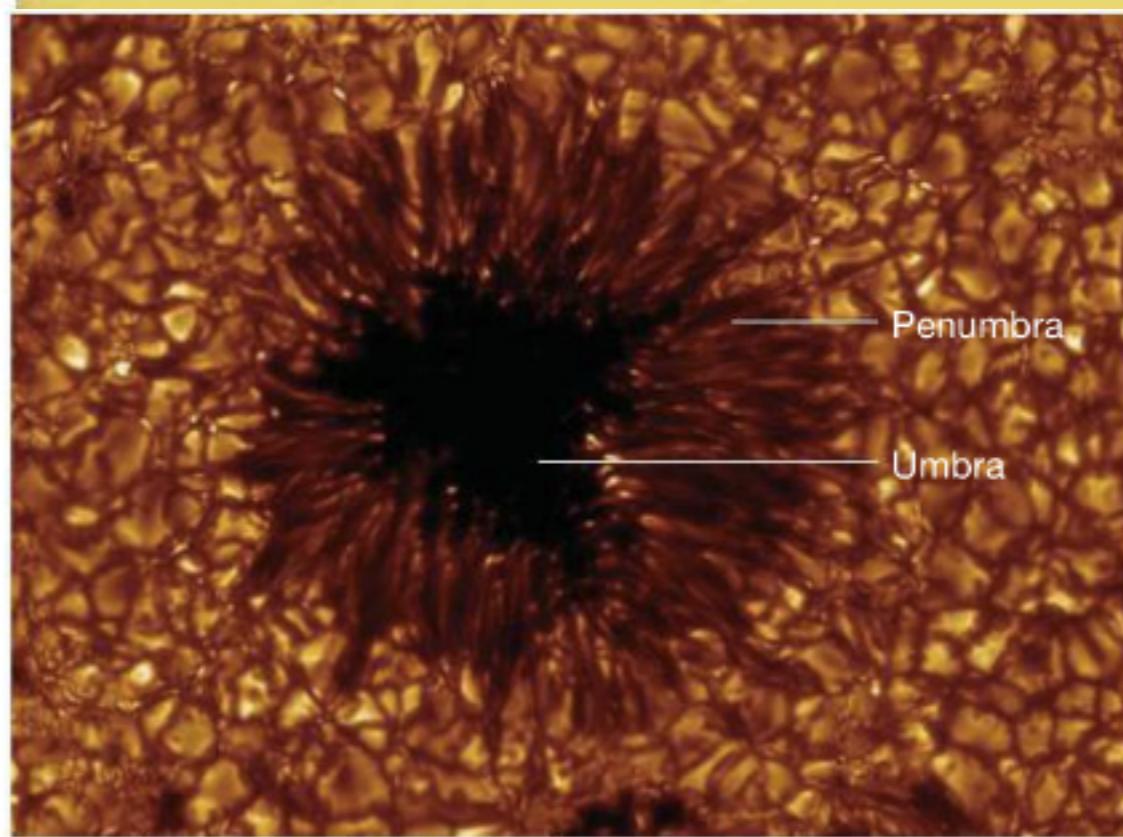
The Sun: Ch 16



Sun Spots

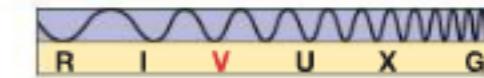


Sunspots appear dark because they are slightly cooler than the surrounding gas.



(b)

← 10,000 km →

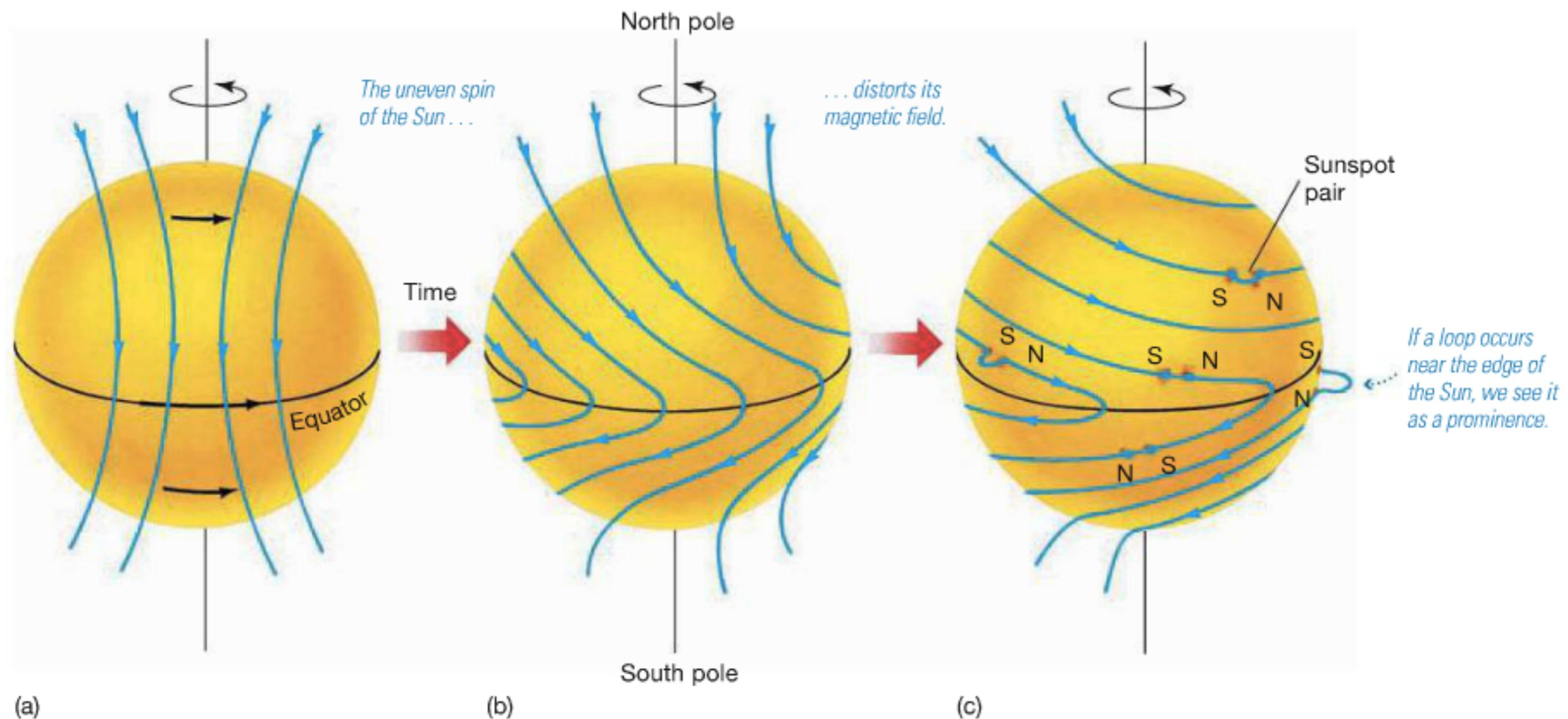


- Umbra is 1000K cooler than the rest of the sun (still bright!)

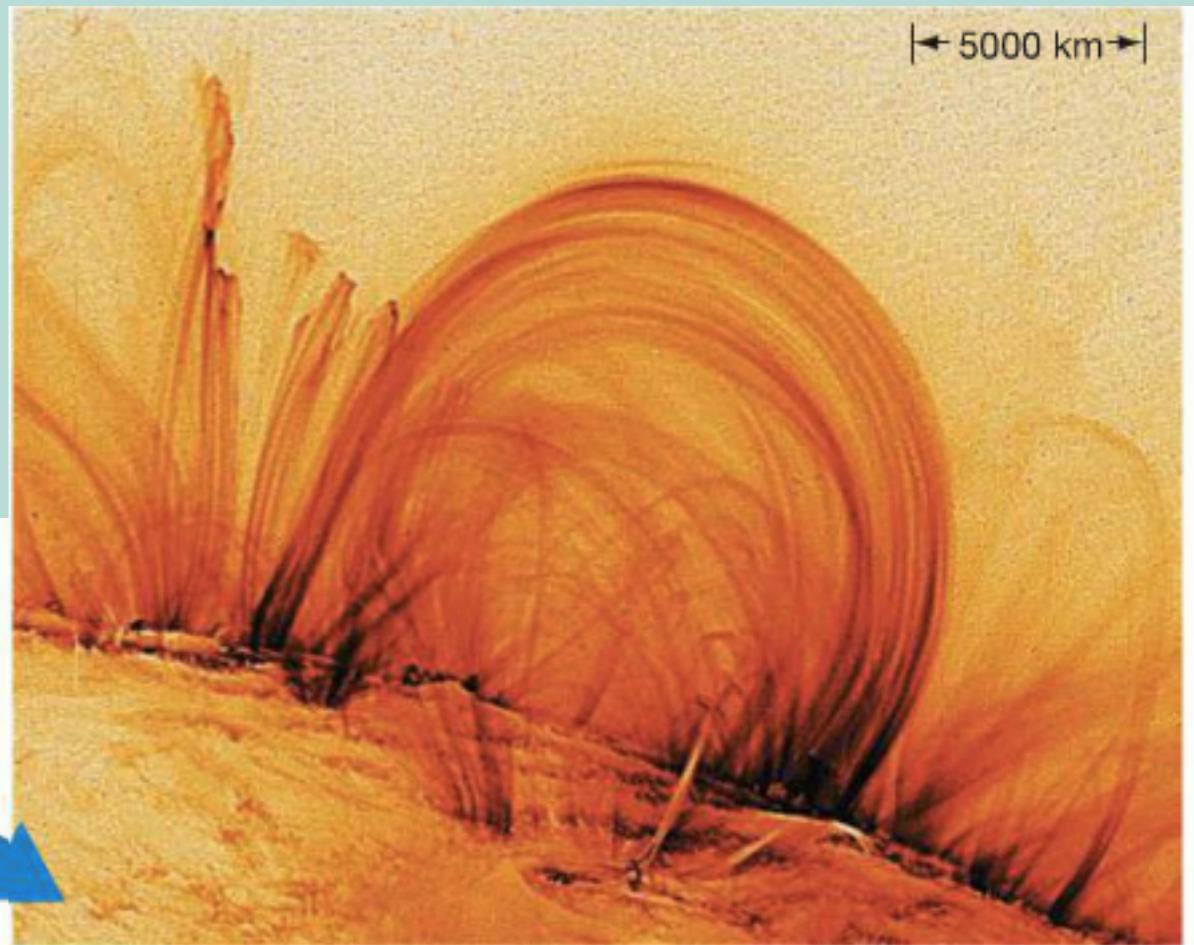
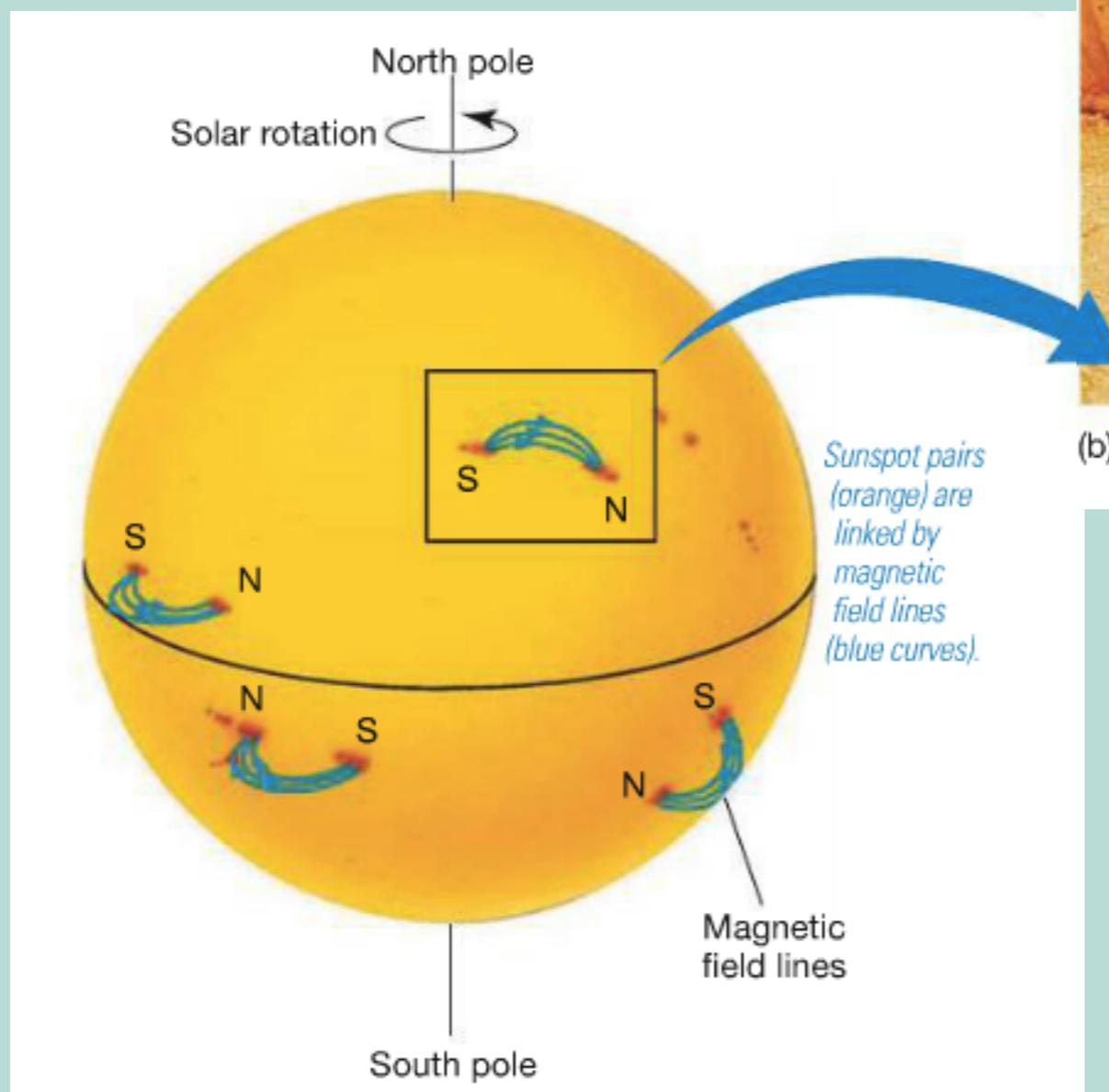


DISCUSSION

Origin of Sun Spots

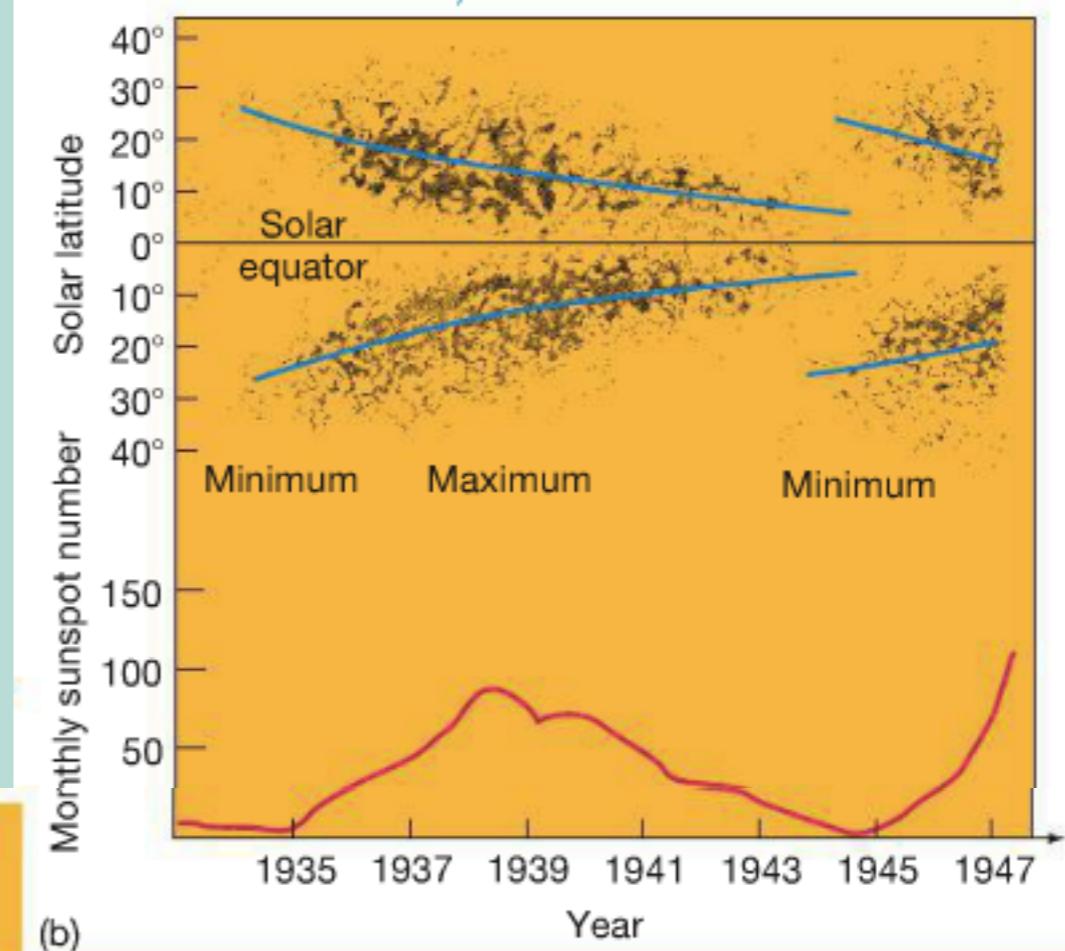
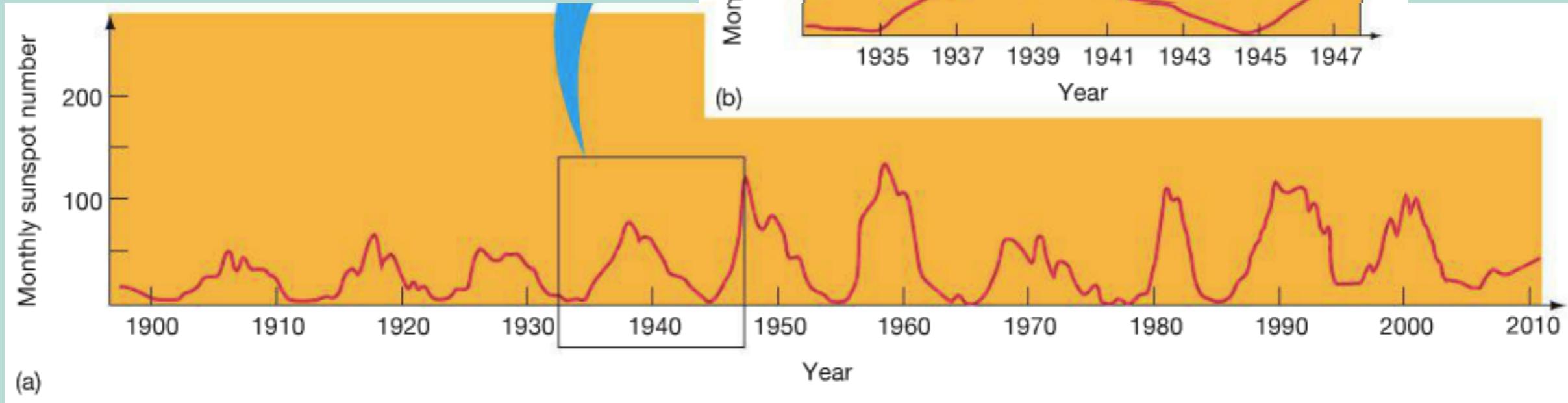


Sun Spots and Magnetism

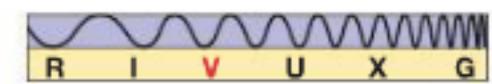
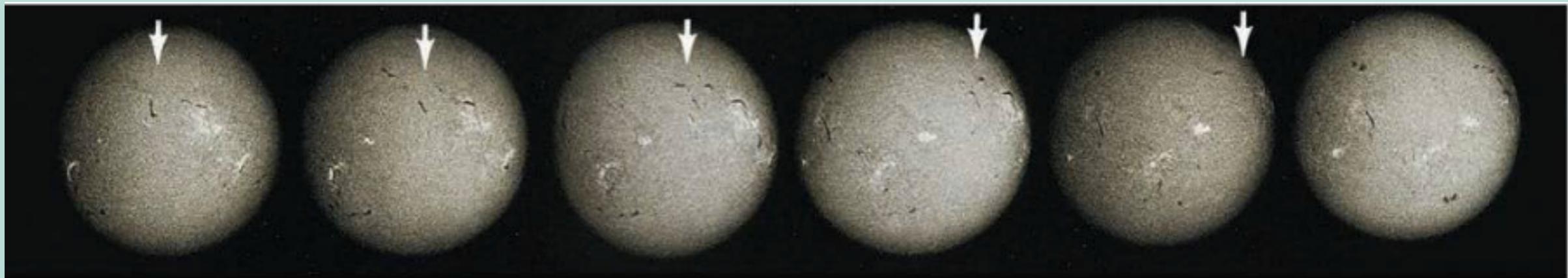
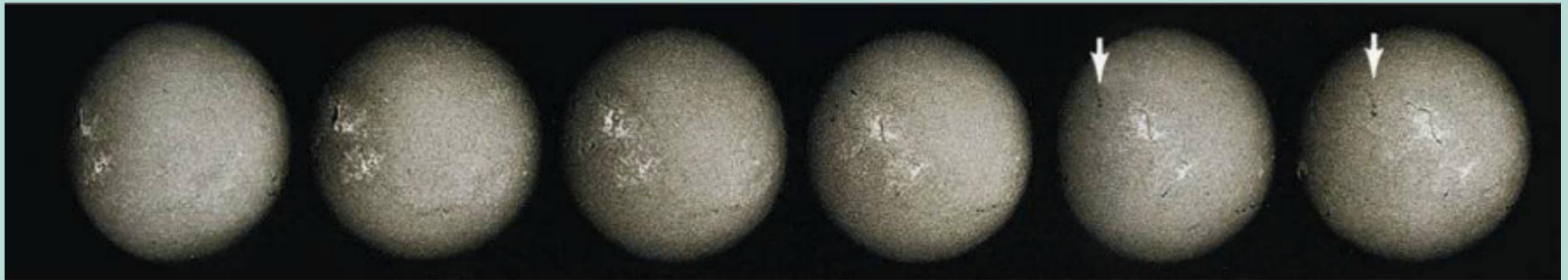


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Solar Cycle:
11 + 11 years

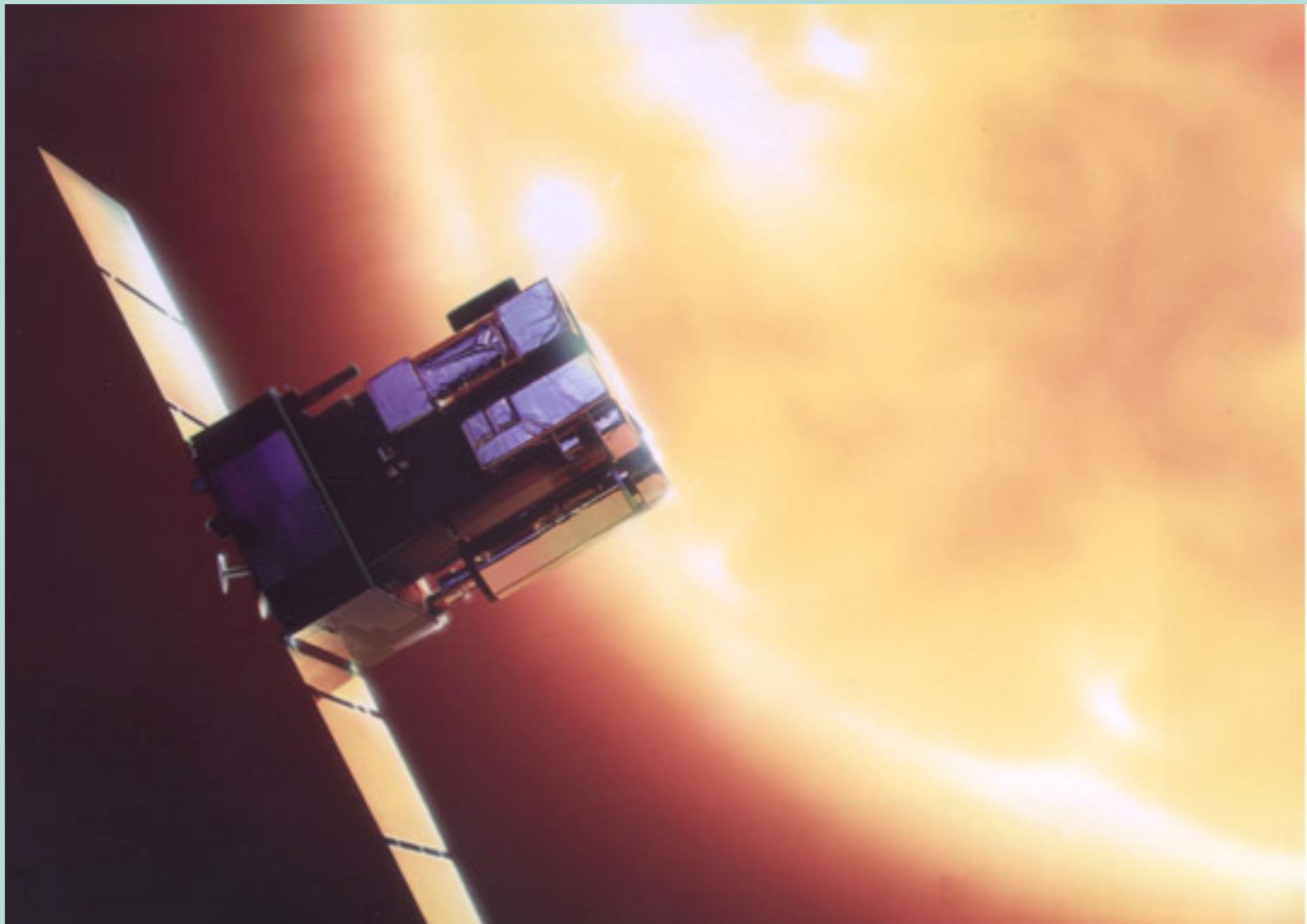


The Sun: Ch 16



Sun Spots in Motion

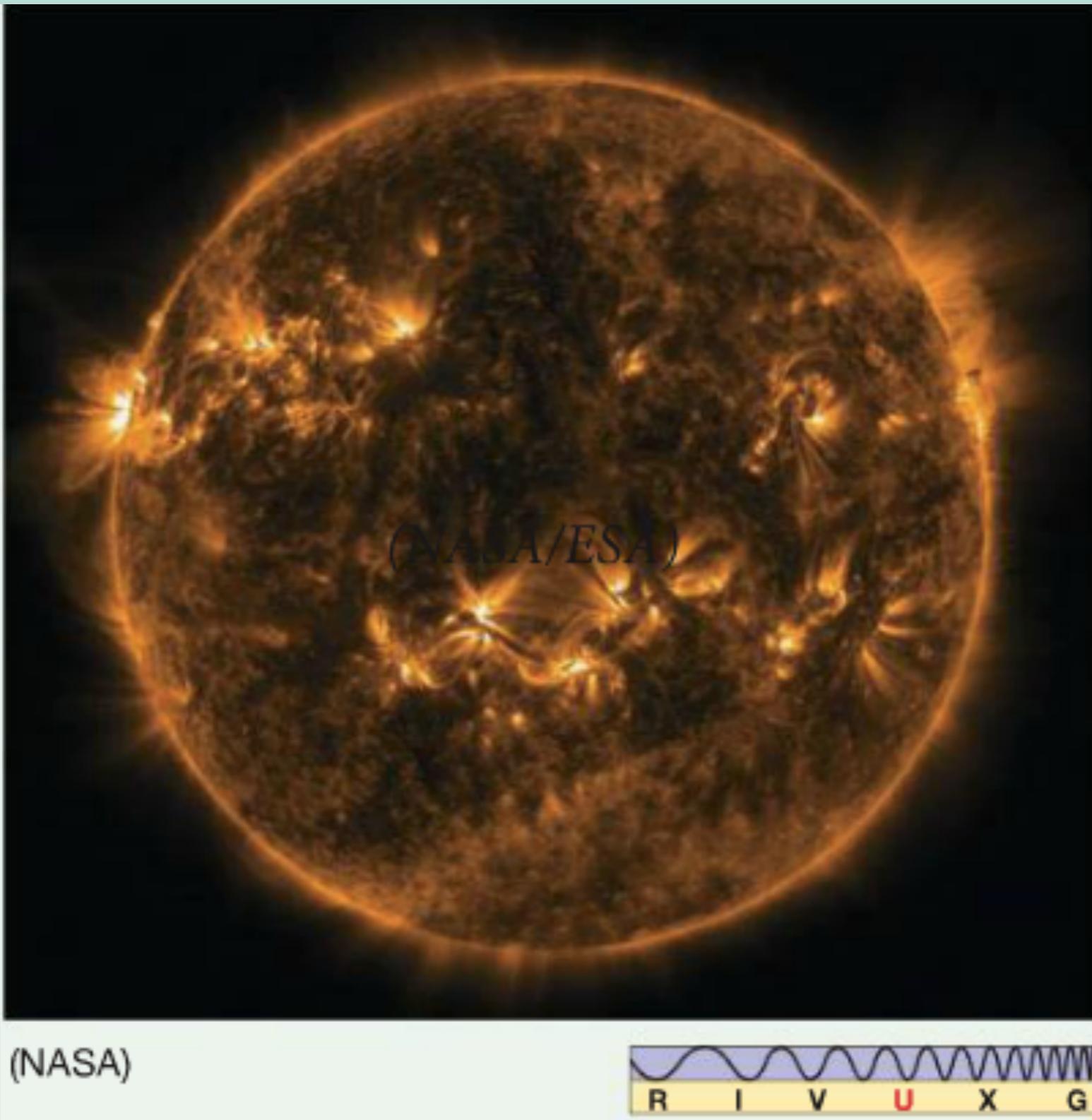
SOHO (SOlar Heliospheric Observatory)



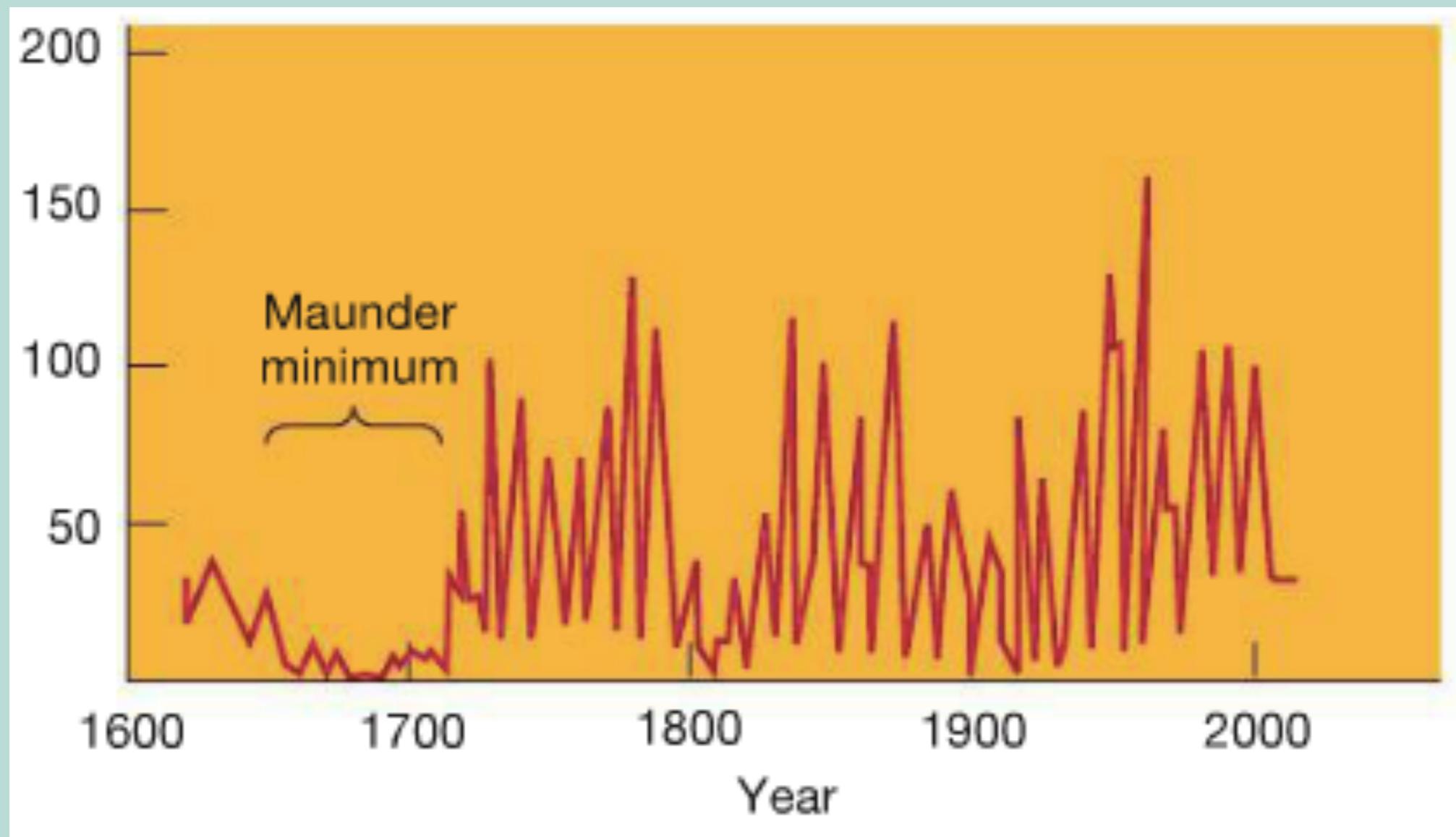
(L1 point)

<https://sohowww.nascom.nasa.gov/>

SDO (Solar Dynamics Observatory)



Maunder Minimum



The Sun: Ch 16

A beautiful summer day in the Netherlands...

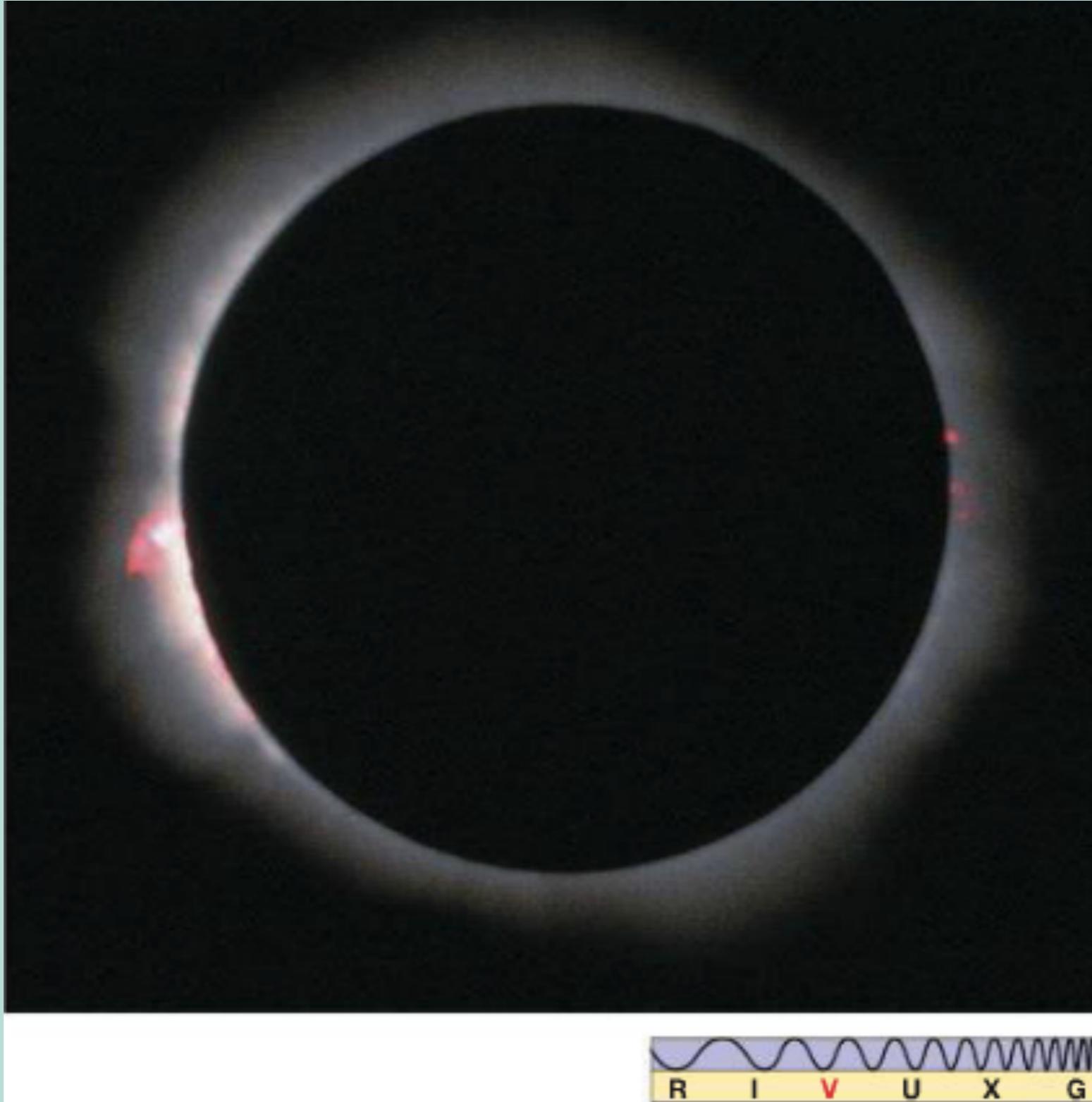


No real evidence of weather



Solar Activity

Chromosphere



(red hue due to Ha lines)

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Corona

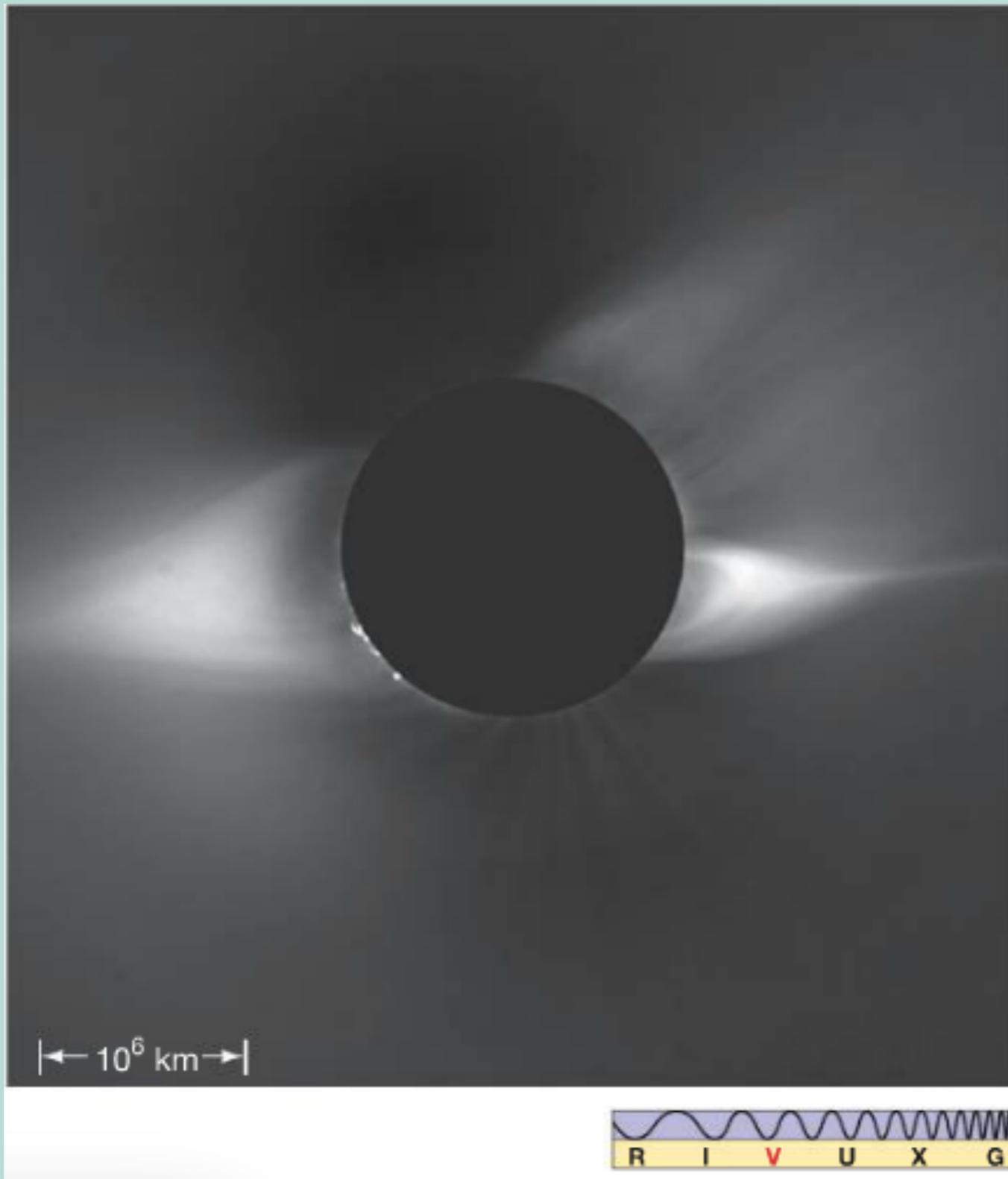


The Sun: Ch 16

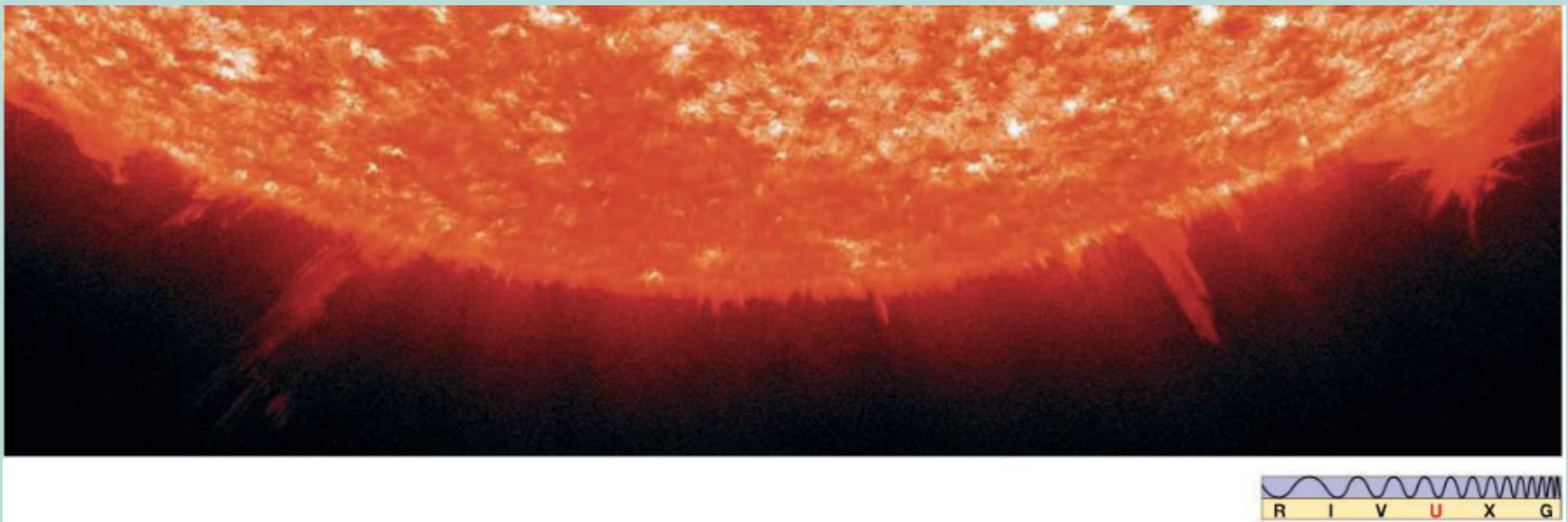
- Solar wind blasts out towards earth at 500 km/s
- Thus the sun is “evaporating” at a rate of 2 million tons of solar matter per second
- Sounds big, but amounts to 0.1% of its mass over the last 4.6 billion years

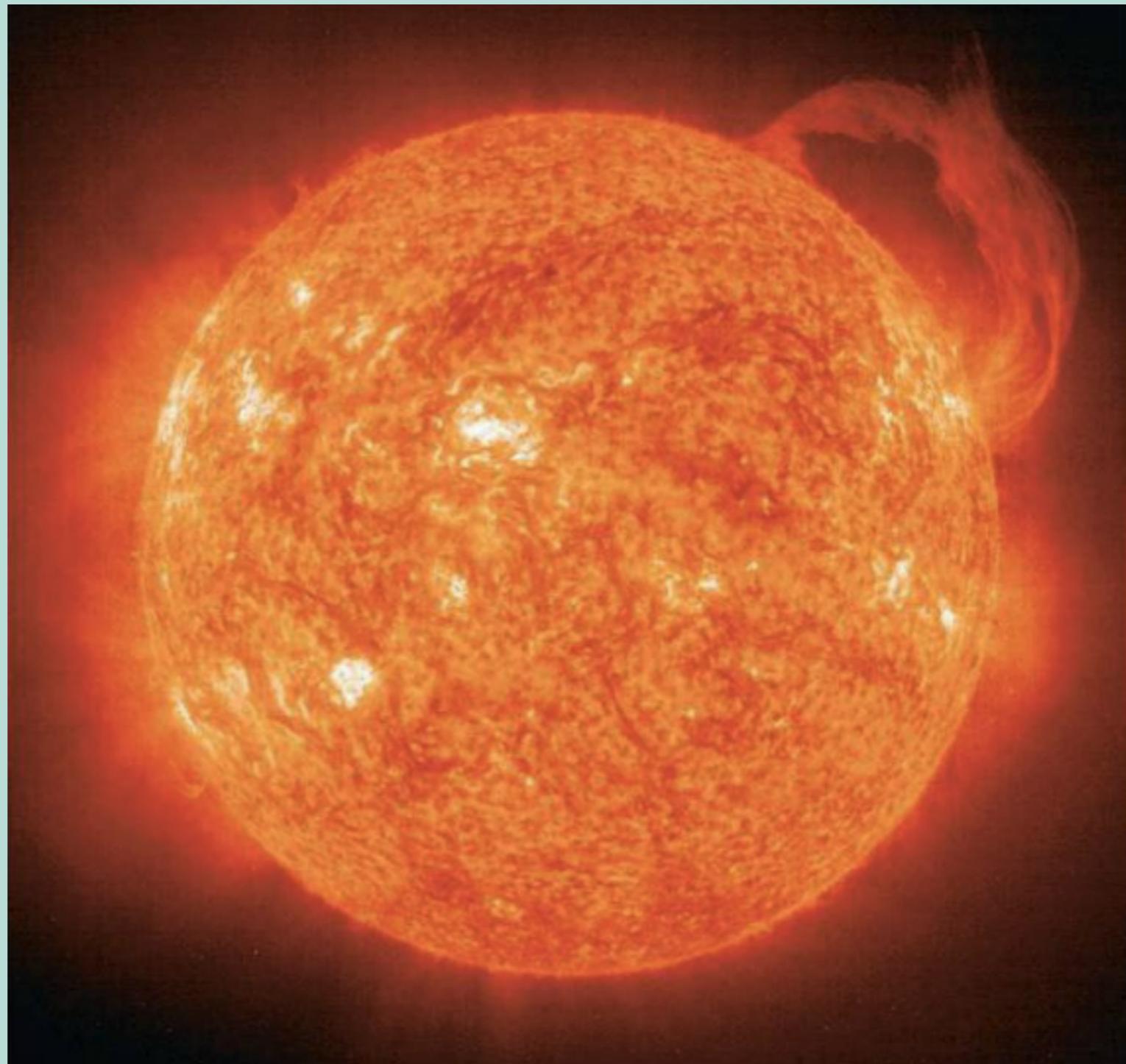
The Sun: Ch 16

Corona during solar cycle peak



Solar Spicules



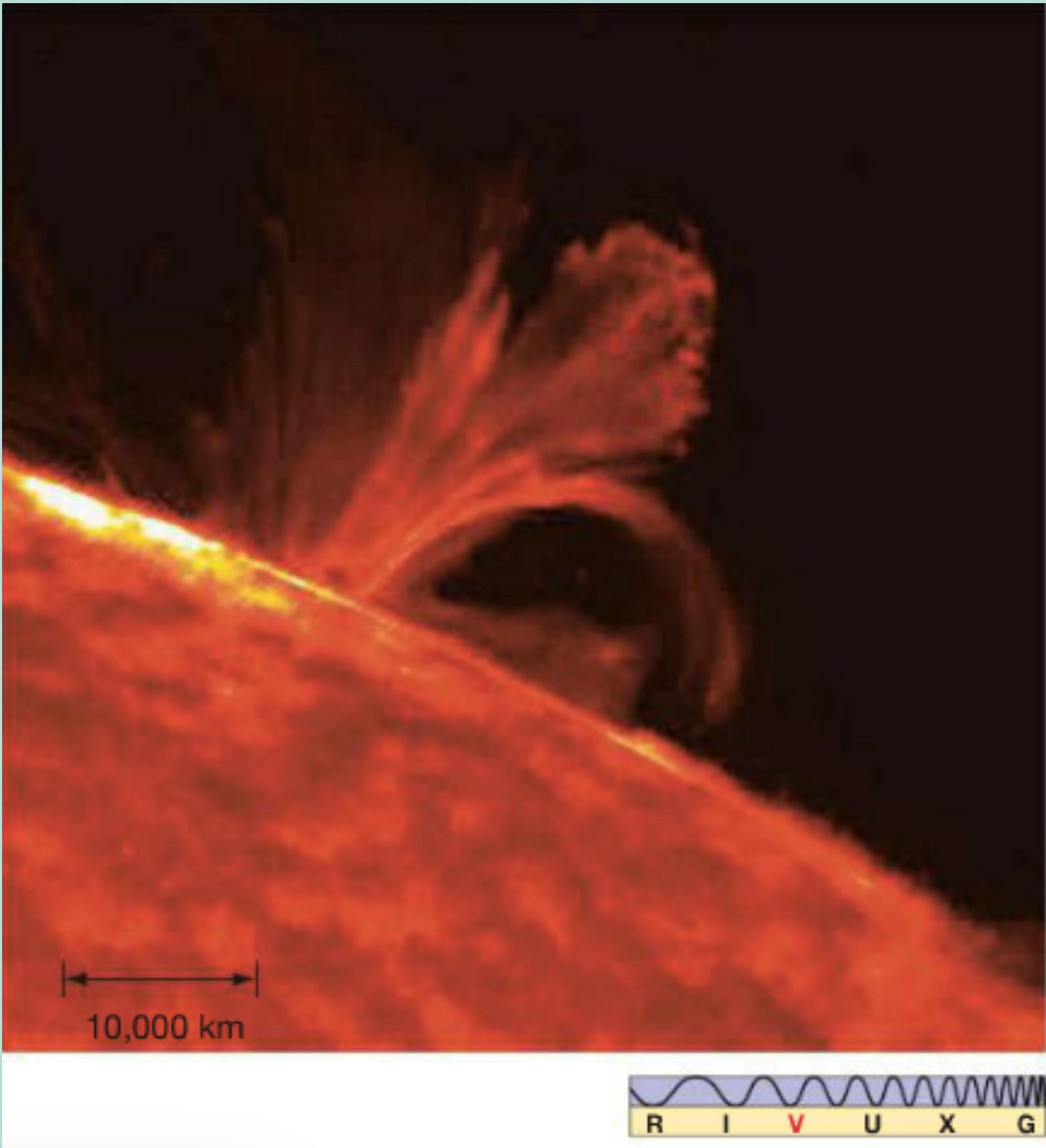


Solar Prominence



The Sun: Ch 16

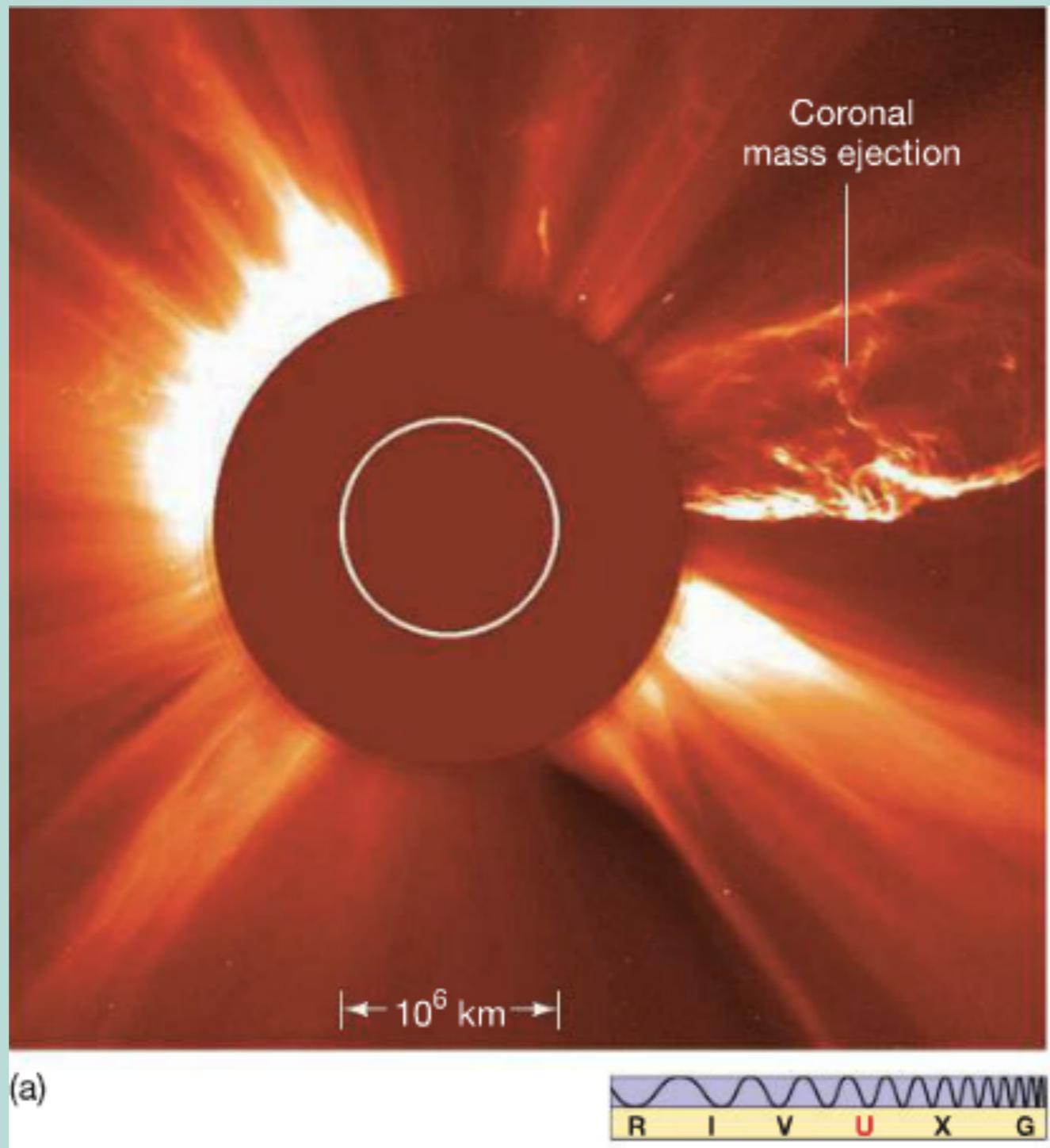
Solar Flare



The “Carrington” Event

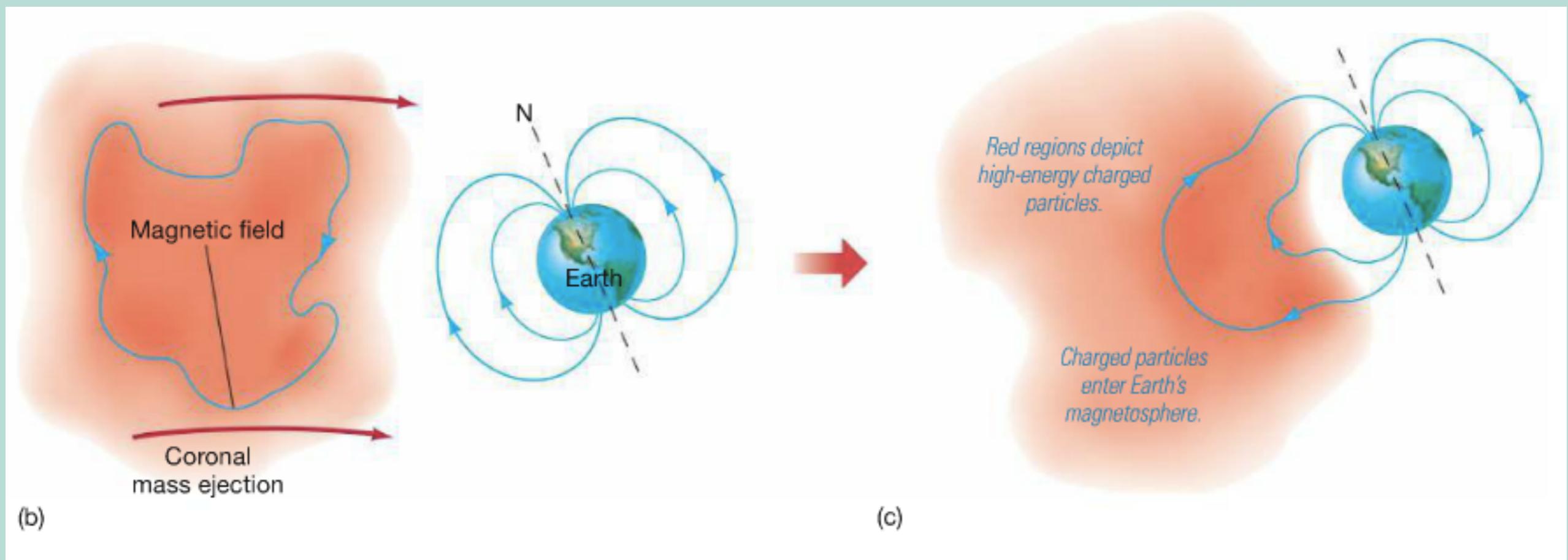
- September 1859
- Aurora could be seen at latitudes as far south as the Caribbean
- The Auroras in the Rockies were so bright they awoke gold miners who began preparing breakfast
- Telegraph operators could still send signals through disconnected lines!
- It has been estimated that if this happened today it would do ~ 2 trillion dollars of damage

Coronal Mass Ejection (CME)

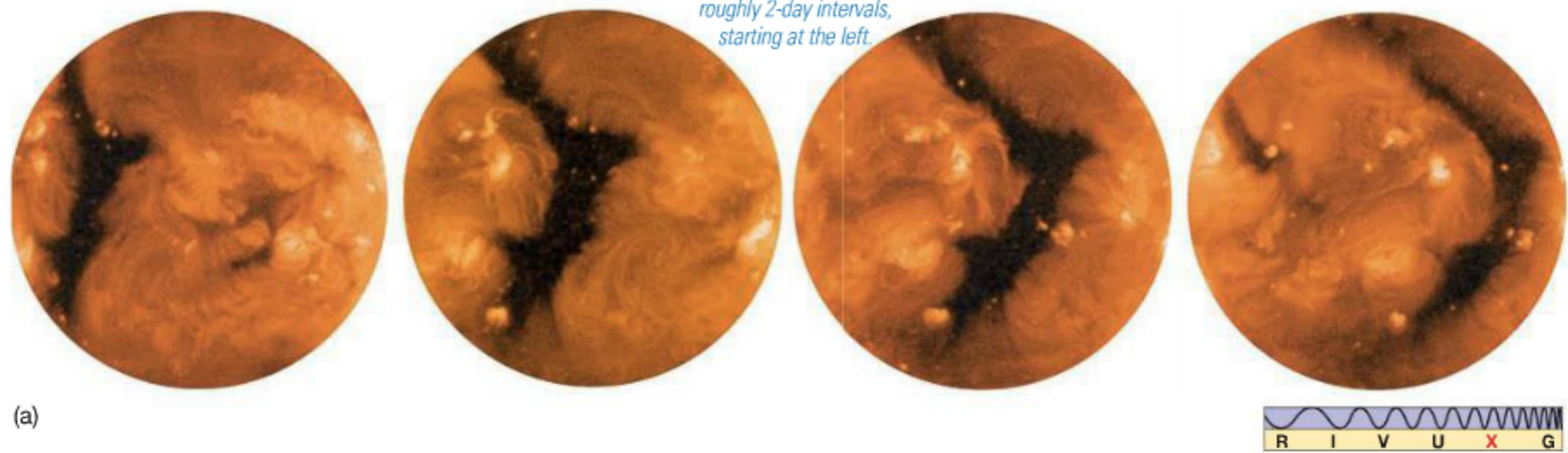


The Sun: Ch 16

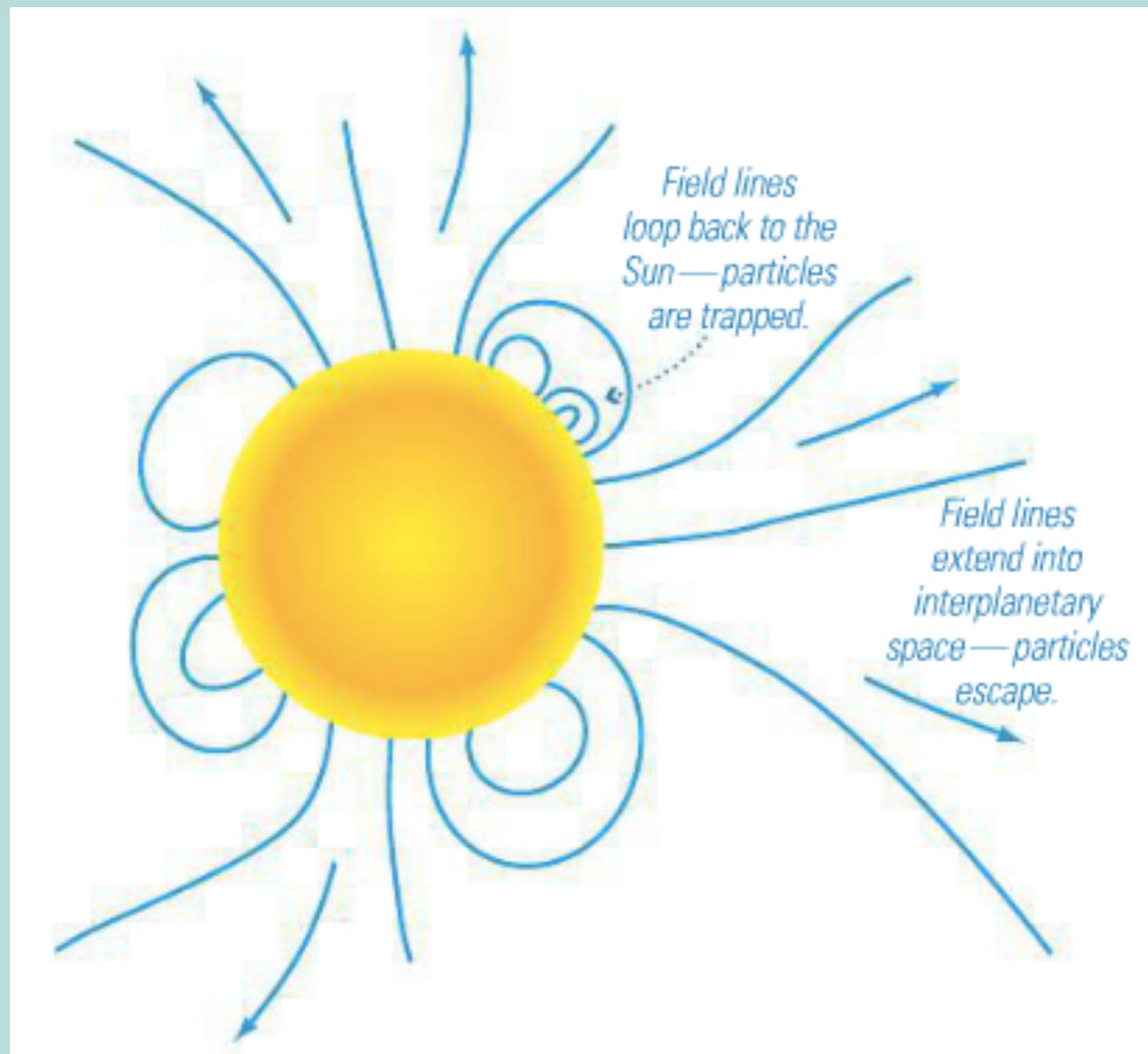
Reconnection



Coronal Holes



“Coronal Holes” are escaping field lines



The Sun: Ch 16

Sir Arthur Eddington



1882 - 1944 UK

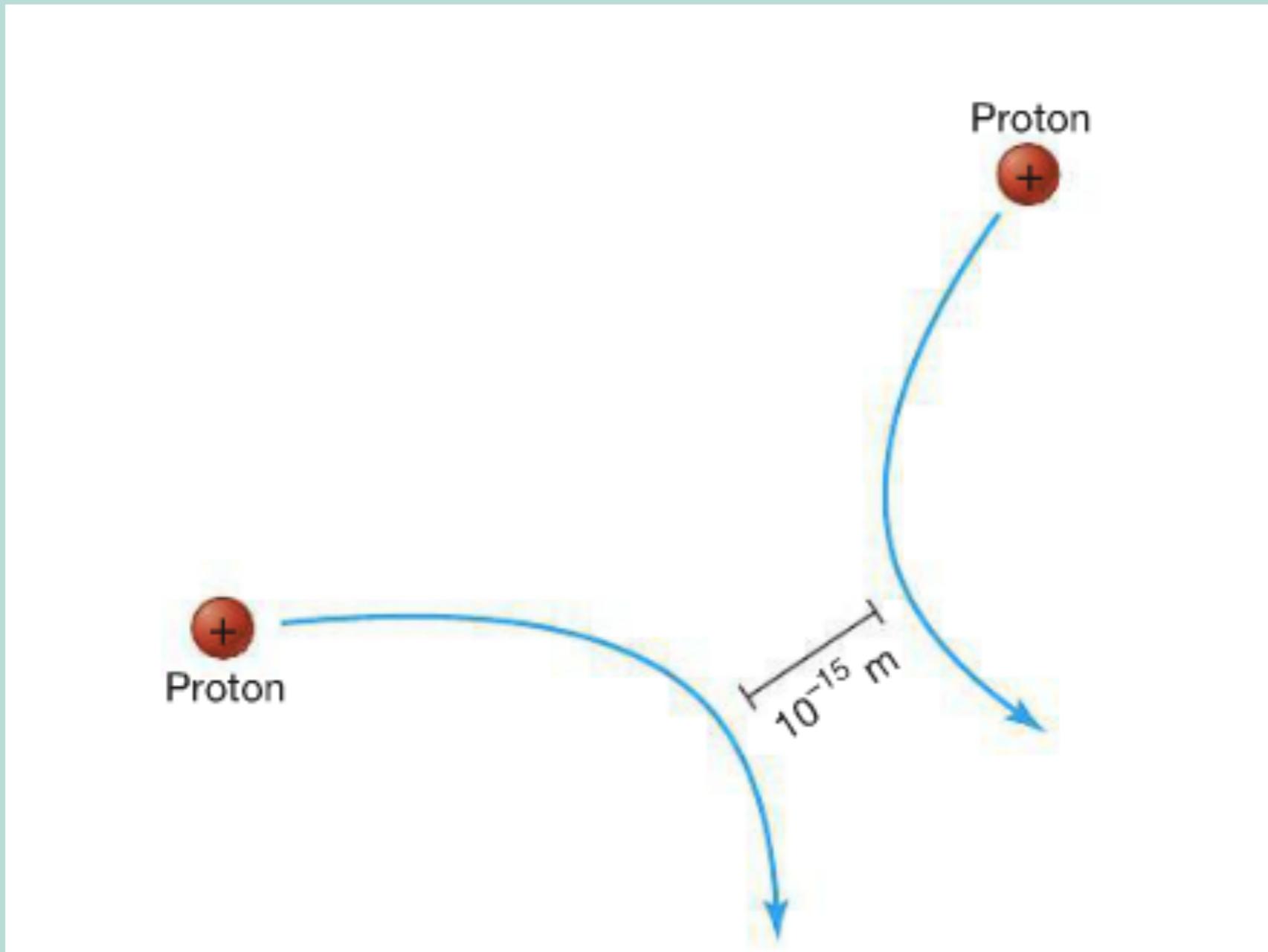
Fusion Vs Wood

- If you divide the solar luminosity by the solar mass you get 2 mW / kilogram
- Wood generates 1,000,000 times more than that. A bon fire, merely the size of the pitiful earth would burn as bright as the sun
- How long would that fire burn though? Ans: not 5 billion years.
- Consider over 5 billion years, the amount of energy each kilogram of sun has provided so far:

10^{13} Joules!!

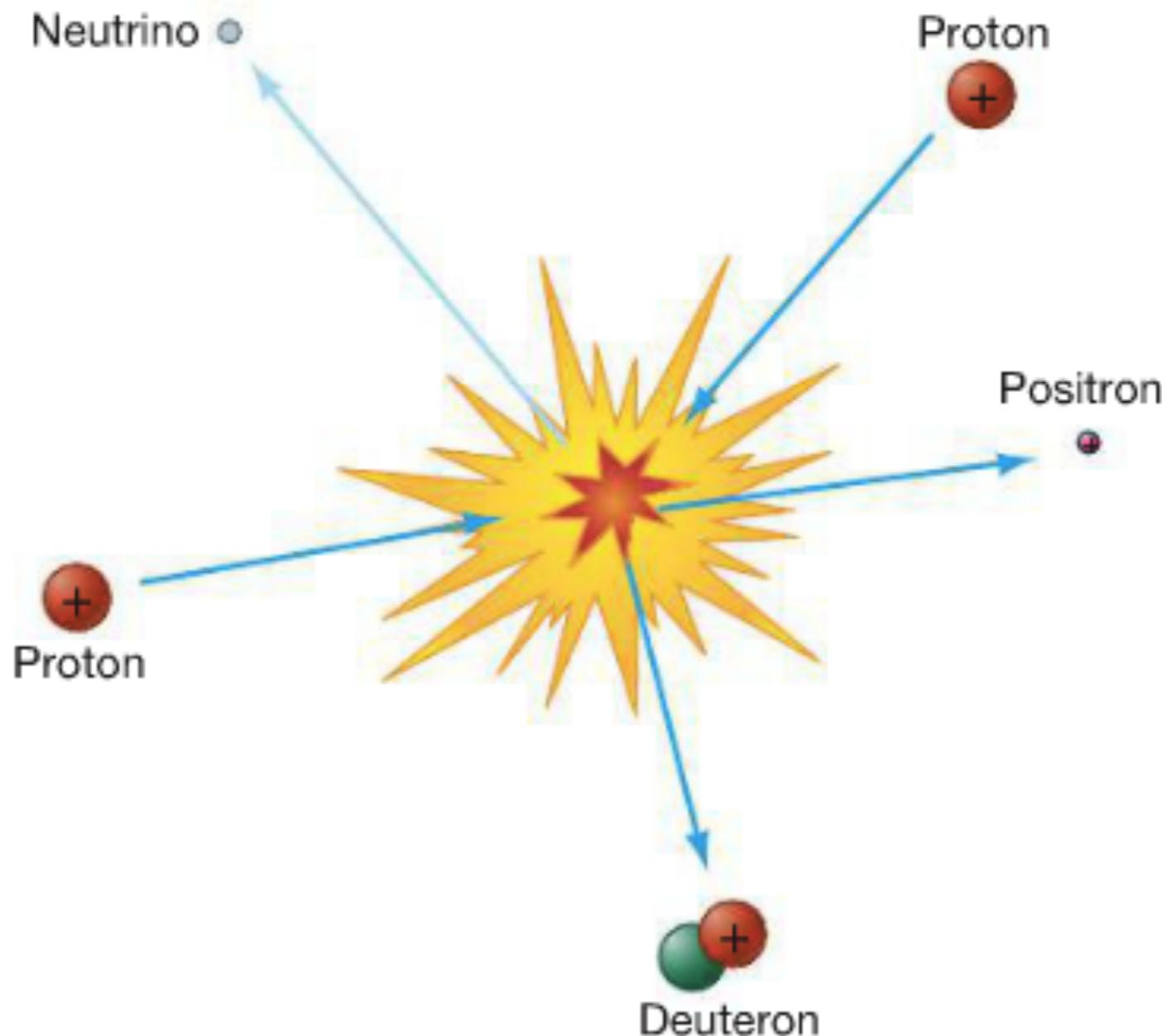
$$E = mc^2$$

- This means 1 kg in principle could have 10^{17} Joules packed inside!



Electrostatic Repulsion

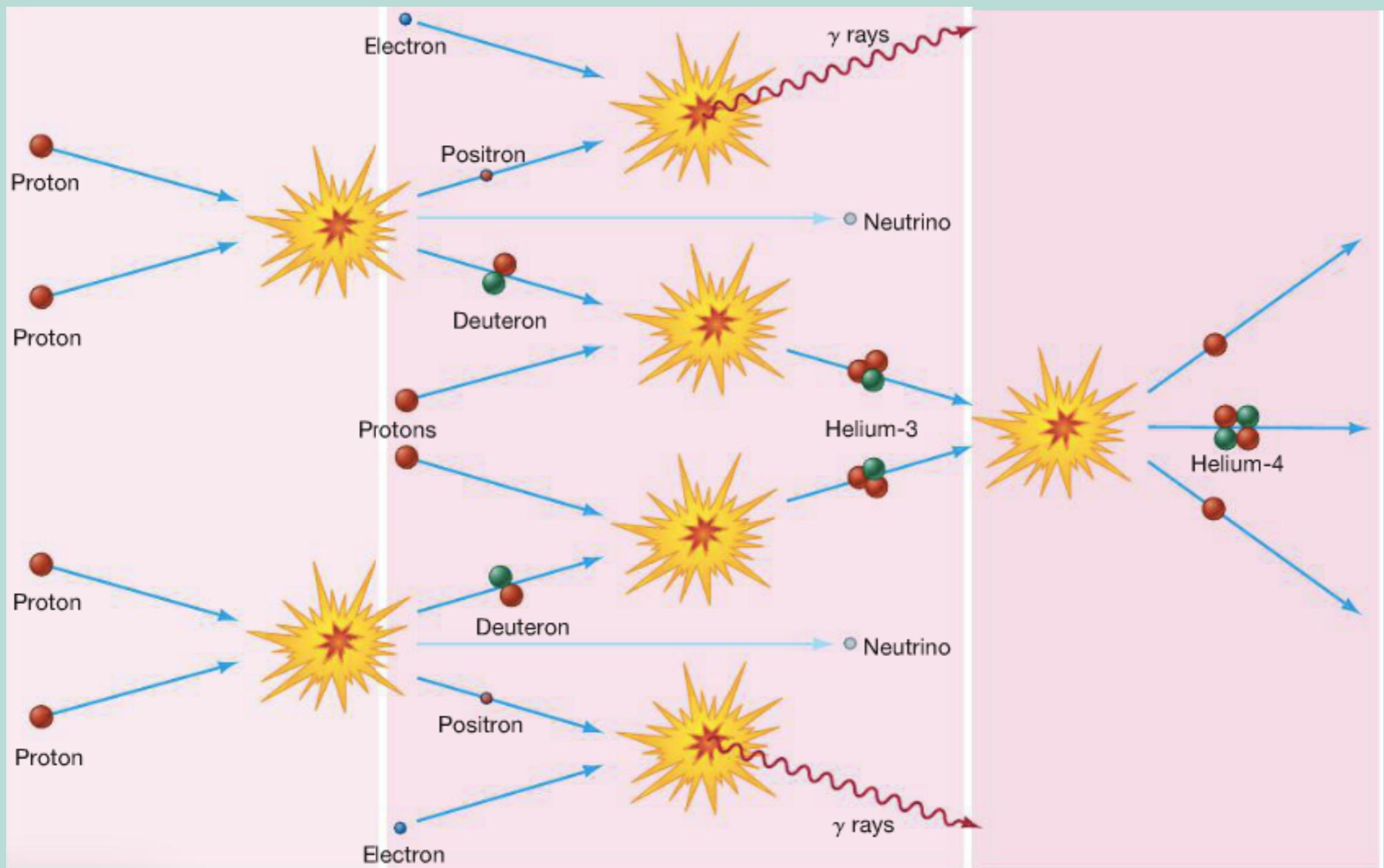
(we'll talk about these later)

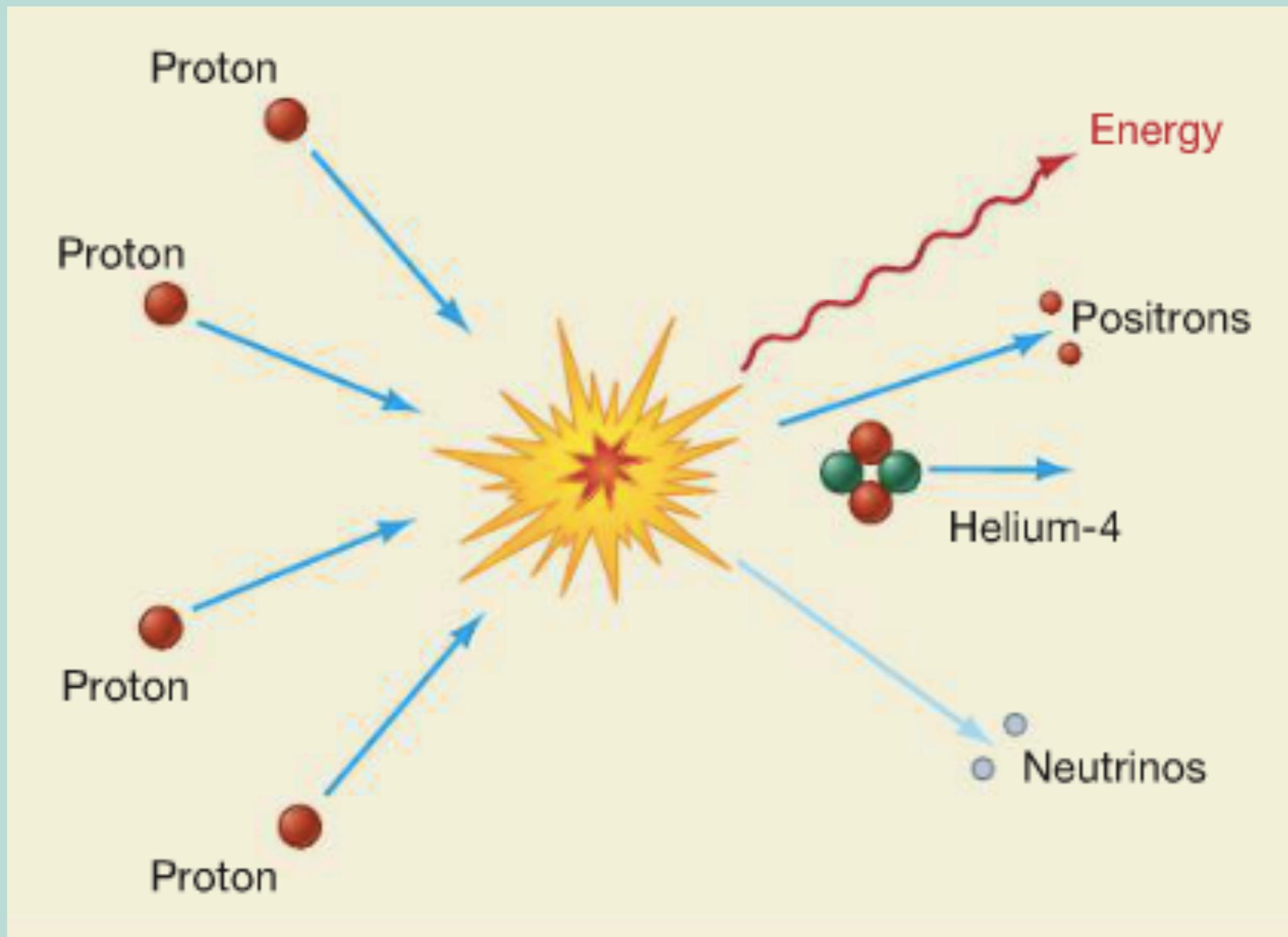


(anti-matter)

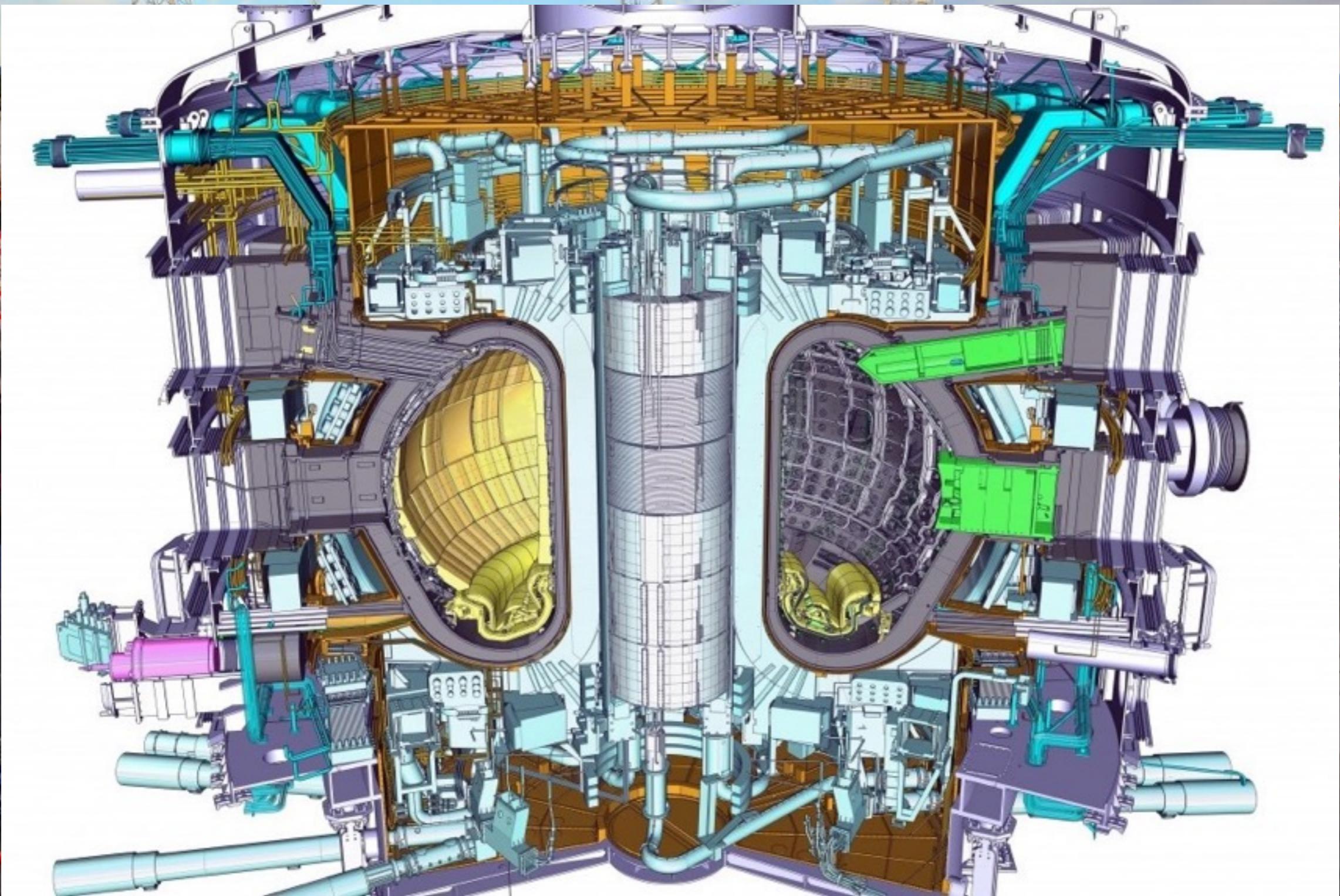
(isotopes)

How to burn 4 hydrogen nuclei





- Each fusion reaction gives you 4×10^{-12} J of energy.
- 1 kg gives you 10^{14} J of energy, so fusion (as a fraction of the total mass energy) has an efficiency of .71%
- Side note, dropping rocks into maximally rotating black holes has an efficiency of 47%
- Thus, to keep up with 4×10^{26} W of luminosity, 600 million tons of hydrogen must be fused into helium every second.
- Thus we lose 4.3 million tons of mass to pure energy per second (comparable to the solar wind).
- That might sound like a lot, but the sun is 10^{30} kg, so we can keep doing this for 5 billion years



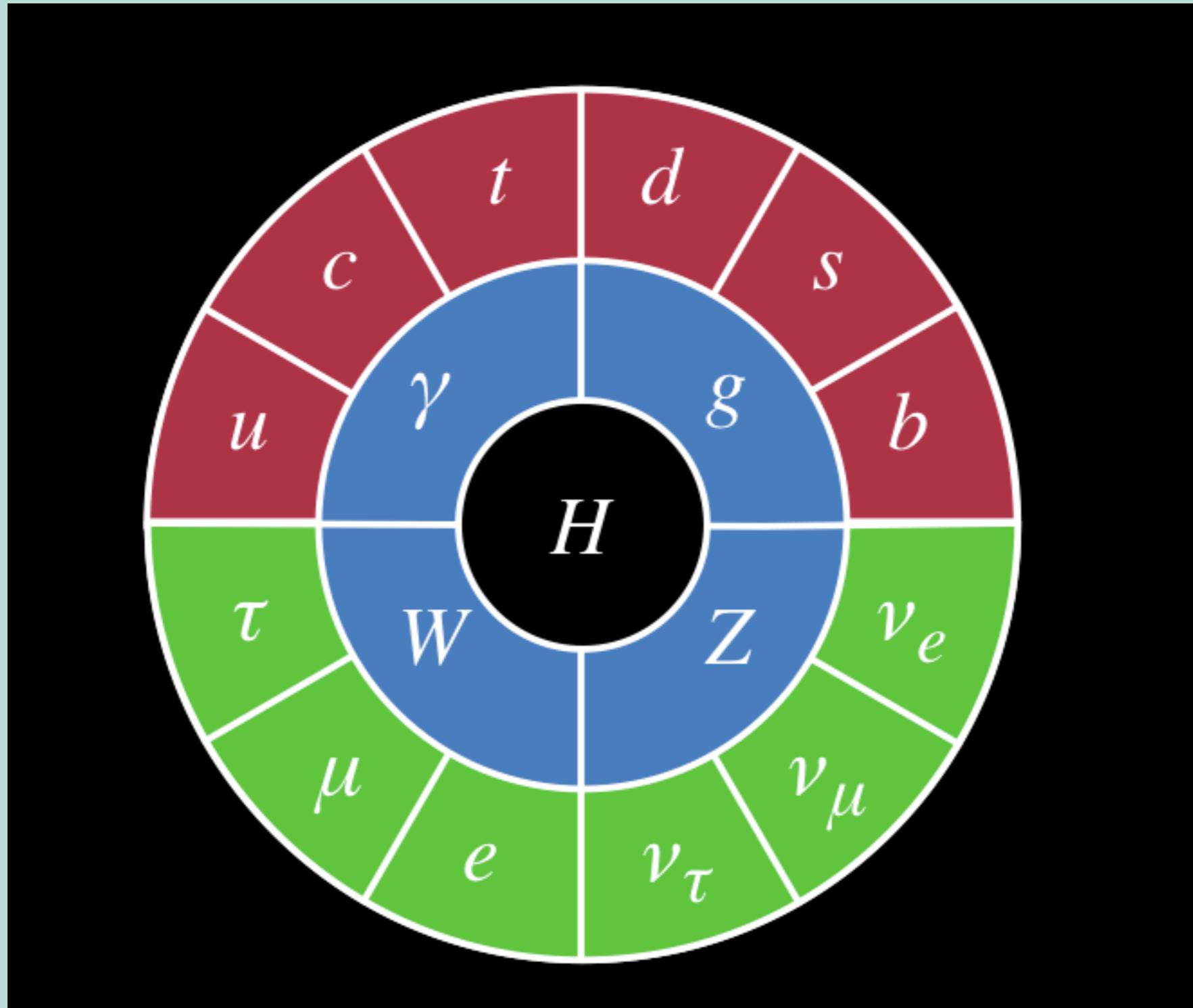
International ThermoNuclear Reactor (2027)



DISCUSSION

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Behold...



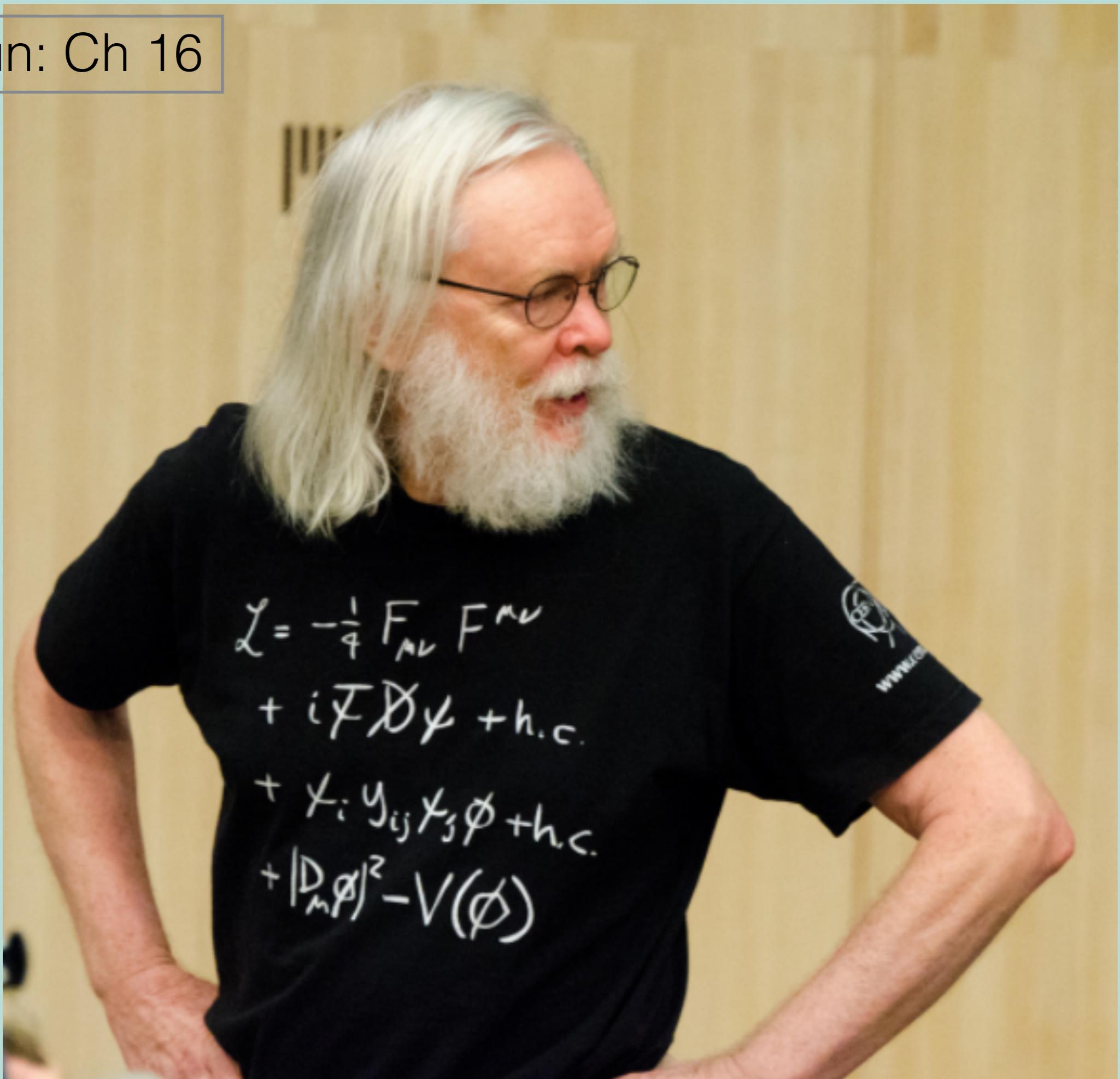


The Sun: Ch 16

75% of Physics.

$$\begin{aligned}
\mathcal{L}_{SM} = & -\frac{1}{2}\partial_\nu g_\mu^a \partial_\nu g_\mu^a - g_s f^{abc} \partial_\mu g_\nu^a g_\mu^b g_\nu^c - \frac{1}{4}g_e^2 f^{abc} f^{ade} g_\mu^b g_\nu^c g_\mu^d g_\nu^e - \partial_\nu W_\mu^+ \partial_\nu W_\mu^- - \\
& M^2 W_\mu^+ W_\mu^- - \frac{1}{2}\partial_\nu Z_\mu^0 \partial_\nu Z_\mu^0 - \frac{1}{2c_w^2} M^2 Z_\mu^0 Z_\mu^0 - \frac{1}{2}\partial_\mu A_\nu \partial_\mu A_\nu - ig c_w (\partial_\nu Z_\mu^0 (W_\mu^+ W_\nu^- - \\
& W_\nu^+ W_\mu^-) - Z_\nu^0 (W_\mu^+ \partial_\nu W_\mu^- - W_\mu^- \partial_\nu W_\mu^+) + Z_\mu^0 (W_\nu^+ \partial_\nu W_\mu^- - W_\nu^- \partial_\nu W_\mu^+)) - \\
& igs_w (\partial_\nu A_\mu (W_\mu^+ W_\nu^- - W_\nu^+ W_\mu^-) - A_\nu (W_\mu^+ \partial_\nu W_\mu^- - W_\mu^- \partial_\nu W_\mu^+) + A_\mu (W_\nu^+ \partial_\nu W_\mu^- - \\
& W_\nu^- \partial_\nu W_\mu^+)) - \frac{1}{2}g^2 W_\mu^+ W_\mu^- W_\nu^+ W_\nu^- + \frac{1}{2}g^2 W_\mu^+ W_\nu^- W_\nu^+ W_\mu^- + g^2 c_w^2 (Z_\mu^0 W_\mu^+ Z_\nu^0 W_\nu^- - \\
& Z_\mu^0 Z_\mu^0 W_\nu^+ W_\nu^-) + g^2 s_w^2 (A_\mu W_\mu^+ A_\nu W_\nu^- - A_\mu A_\nu W_\nu^+ W_\nu^-) + g^2 s_w c_w (A_\mu Z_\nu^0 (W_\mu^+ W_\nu^- - \\
& W_\nu^+ W_\mu^-) - 2 A_\mu Z_\mu^0 W_\nu^+ W_\nu^-) - \frac{1}{2}\partial_\mu H \partial_\mu H - 2M^2 \alpha_h H^2 - \partial_\mu \phi^+ \partial_\mu \phi^- - \frac{1}{2}\partial_\mu \phi^0 \partial_\mu \phi^0 - \\
& \beta_h \left(\frac{2M^2}{g^2} + \frac{2M}{g} H + \frac{1}{2}(H^2 + \phi^0 \phi^0 + 2\phi^+ \phi^-) \right) + \frac{2M^2}{g^2} \alpha_h - \\
& g \alpha_h M (H^3 + H \phi^0 \phi^0 + 2H \phi^+ \phi^-) - \\
& \frac{1}{8}g^2 \alpha_h (H^4 + (\phi^0)^4 + 4(\phi^+ \phi^-)^2 + 4(\phi^0)^2 \phi^+ \phi^- + 4H^2 \phi^+ \phi^- + 2(\phi^0)^2 H^2) - \\
& g M W_\mu^+ W_\mu^- H - \frac{1}{2}g \frac{M}{c_w^2} Z_\mu^0 Z_\mu^0 H - \\
& \frac{1}{2}ig (W_\mu^+ (\phi^0 \partial_\mu \phi^- - \phi^- \partial_\mu \phi^0) - W_\mu^- (\phi^0 \partial_\mu \phi^+ - \phi^+ \partial_\mu \phi^0)) + \\
& \frac{1}{2}g (W_\mu^+ (H \partial_\mu \phi^- - \phi^- \partial_\mu H) + W_\mu^- (H \partial_\mu \phi^+ - \phi^+ \partial_\mu H)) + \frac{1}{2}g \frac{1}{c_w} (Z_\mu^0 (H \partial_\mu \phi^0 - \phi^0 \partial_\mu H) + \\
& M (\frac{1}{c_w} Z_\mu^0 \partial_\mu \phi^0 + W_\mu^+ \partial_\mu \phi^- + W_\mu^- \partial_\mu \phi^+) - ig \frac{s_w^2}{c_w} M Z_\mu^0 (W_\mu^+ \phi^- - W_\mu^- \phi^+) + igs_w M A_\mu (W_\mu^+ \phi^- - \\
& W_\mu^- \phi^+) - ig \frac{1-2c_w^2}{2c_w} Z_\mu^0 (\phi^+ \partial_\mu \phi^- - \phi^- \partial_\mu \phi^+) + igs_w A_\mu (\phi^+ \partial_\mu \phi^- - \phi^- \partial_\mu \phi^+) - \\
& \frac{1}{4}g^2 W_\mu^+ W_\mu^- (H^2 + (\phi^0)^2 + 2\phi^+ \phi^-) - \frac{1}{8}g^2 \frac{1}{c_w^2} Z_\mu^0 Z_\mu^0 (H^2 + (\phi^0)^2 + 2(2s_w^2 - 1)^2 \phi^+ \phi^-) - \\
& \frac{1}{2}g^2 \frac{s_w^2}{c_w} Z_\mu^0 \phi^0 (W_\mu^+ \phi^- + W_\mu^- \phi^+) - \frac{1}{2}ig^2 \frac{s_w^2}{c_w} Z_\mu^0 H (W_\mu^+ \phi^- - W_\mu^- \phi^+) + \frac{1}{2}g^2 s_w A_\mu \phi^0 (W_\mu^+ \phi^- + \\
& W_\mu^- \phi^+) + \frac{1}{2}ig^2 s_w A_\mu H (W_\mu^+ \phi^- - W_\mu^- \phi^+) - g^2 \frac{s_w}{c_w} (2c_w^2 - 1) Z_\mu^0 A_\mu \phi^+ \phi^- - \\
& g^2 s_w^2 A_\mu A_\mu \phi^+ \phi^- + \frac{1}{2}ig_s \lambda_{ij}^a (q_i^\sigma \gamma^\mu q_j^\tau) g_\mu^a - \bar{e}^\lambda (\gamma \partial + m_e^\lambda) e^\lambda - \bar{\nu}^\lambda (\gamma \partial + m_\nu^\lambda) \nu^\lambda - \bar{u}_j^\lambda (\gamma \partial + \\
& m_u^\lambda) u_j^\lambda - \bar{d}_j^\lambda (\gamma \partial + m_d^\lambda) d_j^\lambda + igs_w A_\mu \left(-(\bar{e}^\lambda \gamma^\mu e^\lambda) + \frac{2}{3}(\bar{u}_j^\lambda \gamma^\mu u_j^\lambda) - \frac{1}{3}(\bar{d}_j^\lambda \gamma^\mu d_j^\lambda) \right) + \\
& \frac{ig}{4c_w} Z_\mu^0 \left\{ (\bar{\nu}^\lambda \gamma^\mu (1 + \gamma^5) \nu^\lambda) + (\bar{e}^\lambda \gamma^\mu (4s_w^2 - 1 - \gamma^5) e^\lambda) + (\bar{d}_j^\lambda \gamma^\mu (\frac{4}{3}s_w^2 - 1 - \gamma^5) d_j^\lambda) + \right. \\
& \left. (\bar{u}_j^\lambda \gamma^\mu (1 - \frac{8}{3}s_w^2 + \gamma^5) u_j^\lambda) \right\} + \frac{ig}{2\sqrt{2}} W_\mu^+ \left((\bar{\nu}^\lambda \gamma^\mu (1 + \gamma^5) U^{lep}{}^\lambda{}_\kappa e^\kappa) + (\bar{u}_j^\lambda \gamma^\mu (1 + \gamma^5) C_{\lambda\kappa} d_j^\kappa) \right) + \\
& \frac{ig}{2\sqrt{2}} W_\mu^- \left((\bar{e}^\kappa U^{lep\dagger}{}_{\kappa\lambda} \gamma^\mu (1 + \gamma^5) \nu^\lambda) + (\bar{d}_j^\lambda C_{\lambda\kappa}^\dagger \gamma^\mu (1 + \gamma^5) u_j^\kappa) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^+ \left(-m_e^\kappa (\bar{\nu}^\lambda U^{lep}{}_{\lambda\kappa} (1 - \gamma^5) e^\kappa) + m_\nu^\lambda (\bar{\nu}^\lambda U^{lep}{}_{\lambda\kappa} (1 + \gamma^5) e^\kappa) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^- \left(m_e^\lambda (\bar{e}^\lambda U^{lep\dagger}{}_{\lambda\kappa} (1 + \gamma^5) \nu^\kappa) - m_\nu^\kappa (\bar{e}^\lambda U^{lep\dagger}{}_{\lambda\kappa} (1 - \gamma^5) \nu^\kappa) \right) - \frac{g}{2} \frac{m_e^\lambda}{M} H (\bar{\nu}^\lambda \nu^\lambda) - \\
& \frac{g}{2} \frac{m_e^\lambda}{M} H (\bar{e}^\lambda e^\lambda) + \frac{ig}{2} \frac{m_e^\lambda}{M} \phi^0 (\bar{e}^\lambda \gamma^5 \nu^\lambda) - \frac{ig}{2} \frac{m_\nu^\lambda}{M} \phi^0 (\bar{e}^\lambda \gamma^5 e^\lambda) - \frac{1}{4} \bar{\nu}_\lambda M_{\lambda\kappa}^R (1 - \gamma_5) \bar{\nu}_\kappa - \\
& \frac{1}{4} \bar{\nu}_\lambda M_{\lambda\kappa}^R (1 - \gamma_5) \bar{\nu}_\kappa + \frac{ig}{2M\sqrt{2}} \phi^+ \left(-m_d^\kappa (\bar{u}_j^\lambda C_{\lambda\kappa} (1 - \gamma^5) d_j^\kappa) + m_u^\lambda (\bar{u}_j^\lambda C_{\lambda\kappa} (1 + \gamma^5) d_j^\kappa) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^- \left(m_d^\lambda (\bar{d}_j^\lambda C_{\lambda\kappa}^\dagger (1 + \gamma^5) u_j^\kappa) - m_u^\kappa (\bar{d}_j^\lambda C_{\lambda\kappa}^\dagger (1 - \gamma^5) u_j^\kappa) \right) - \frac{g}{2} \frac{m_d^\lambda}{M} H (\bar{u}_j^\lambda u_j^\lambda) - \\
& \frac{g}{2} \frac{m_d^\lambda}{M} H (\bar{d}_j^\lambda d_j^\lambda) + \frac{ig}{2} \frac{m_d^\lambda}{M} \phi^0 (\bar{u}_j^\lambda \gamma^5 u_j^\lambda) - \frac{ig}{2} \frac{m_u^\lambda}{M} \phi^0 (\bar{d}_j^\lambda \gamma^5 d_j^\lambda) + \bar{G}^a \partial^2 G^a + g_s f^{abc} \partial_\mu \bar{G}^a G^b g_\mu^c + \\
& \bar{X}^+ (\partial^2 - M^2) X^+ + \bar{X}^- (\partial^2 - M^2) X^- + \bar{X}^0 (\partial^2 - \frac{M^2}{c_w^2}) X^0 + \bar{Y} \partial^2 Y + ig c_w W_\mu^+ (\partial_\mu \bar{X}^0 X^- - \\
& \partial_\mu \bar{X}^+ X^0) + igs_w W_\mu^+ (\partial_\mu \bar{Y} X^- - \partial_\mu \bar{X}^+ Y) + ig c_w W_\mu^- (\partial_\mu \bar{X}^- X^0 - \\
& \partial_\mu \bar{X}^0 X^+) + igs_w W_\mu^- (\partial_\mu \bar{X}^- Y - \partial_\mu \bar{Y} X^+) + ig c_w Z_\mu^0 (\partial_\mu \bar{X}^+ X^+ - \\
& \partial_\mu \bar{X}^- X^-) + igs_w A_\mu (\partial_\mu \bar{X}^+ X^+ - \\
& \partial_\mu \bar{X}^- X^-) - \frac{1}{2}g M \left(\bar{X}^+ X^+ H + \bar{X}^- X^- H + \frac{1}{c_w^2} \bar{X}^0 X^0 H \right) + \frac{1-2c_w^2}{2c_w} ig M (\bar{X}^+ X^0 \phi^+ - \bar{X}^- X^0 \phi^-) + \\
& \frac{1}{2c_w} ig M (\bar{X}^0 X^- \phi^+ - \bar{X}^0 X^+ \phi^-) + ig M s_w (\bar{X}^0 X^- \phi^+ - \bar{X}^0 X^+ \phi^-) + \\
& \frac{1}{2}ig M (\bar{X}^+ X^+ \phi^0 - \bar{X}^- X^- \phi^0) .
\end{aligned}$$

The Sun: Ch 16



Solar Neutrino Problem

- With the amount of data that we have on the sun, we know exactly how many neutrinos are created in the core
- We also know how likely those neutrinos are to interact with our detectors
- We see a mismatch. This was explained by “neutrino oscillations”, experimentally verified in 2002

The Sun: Ch 16

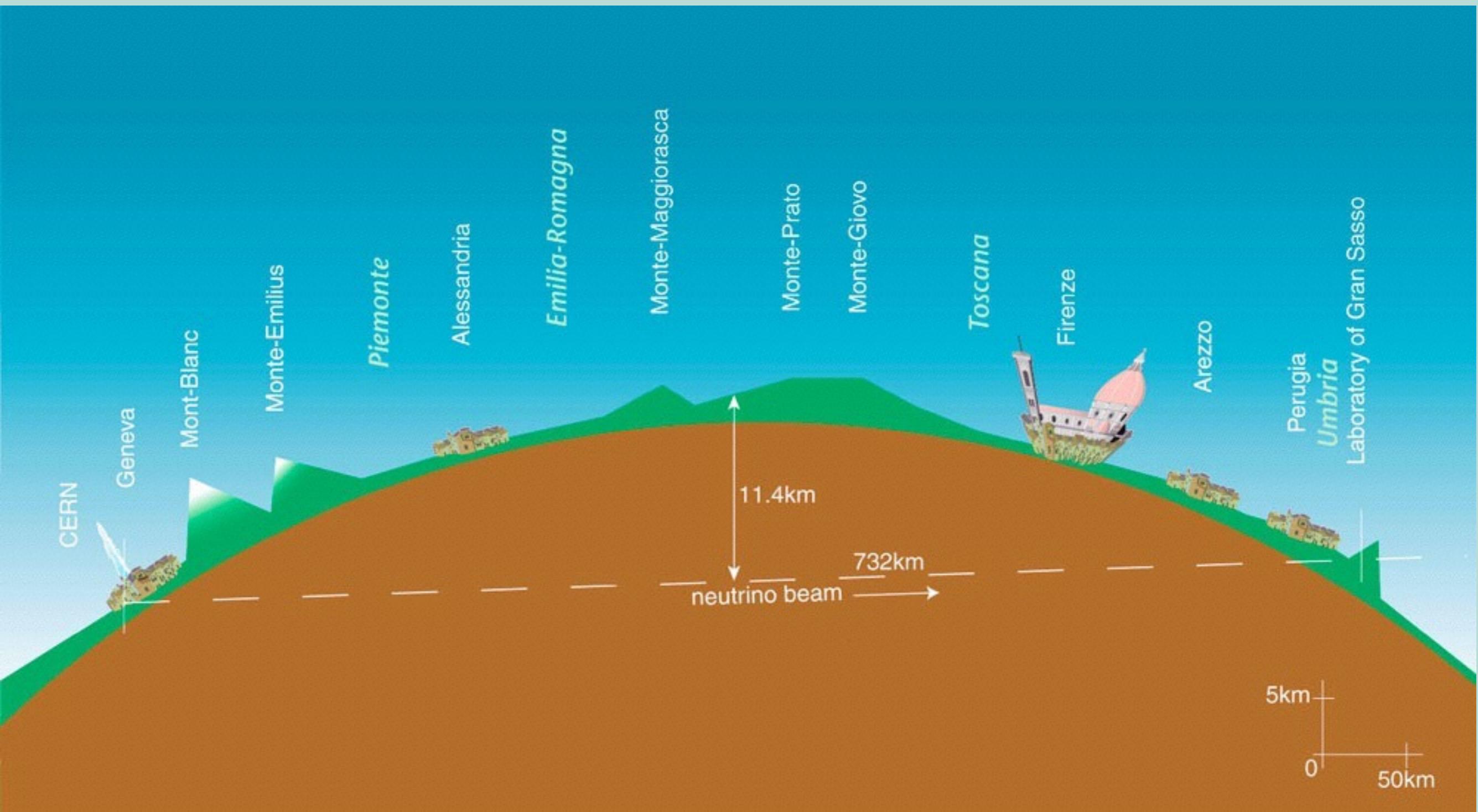


FACT: about 65 million neutrinos pass through your thumbnail every second.

Learn Something
New Every Day
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The Sun: Ch 16

The Earth is transparent to neutrinos.





IceCube Laboratory

Data is collected here and sent by satellite to the data warehouse at UW–Madison



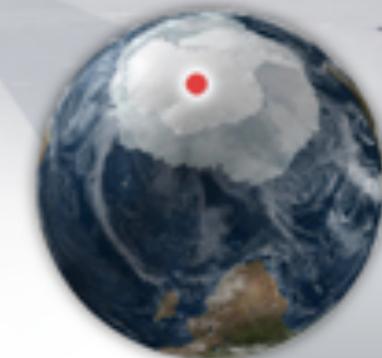
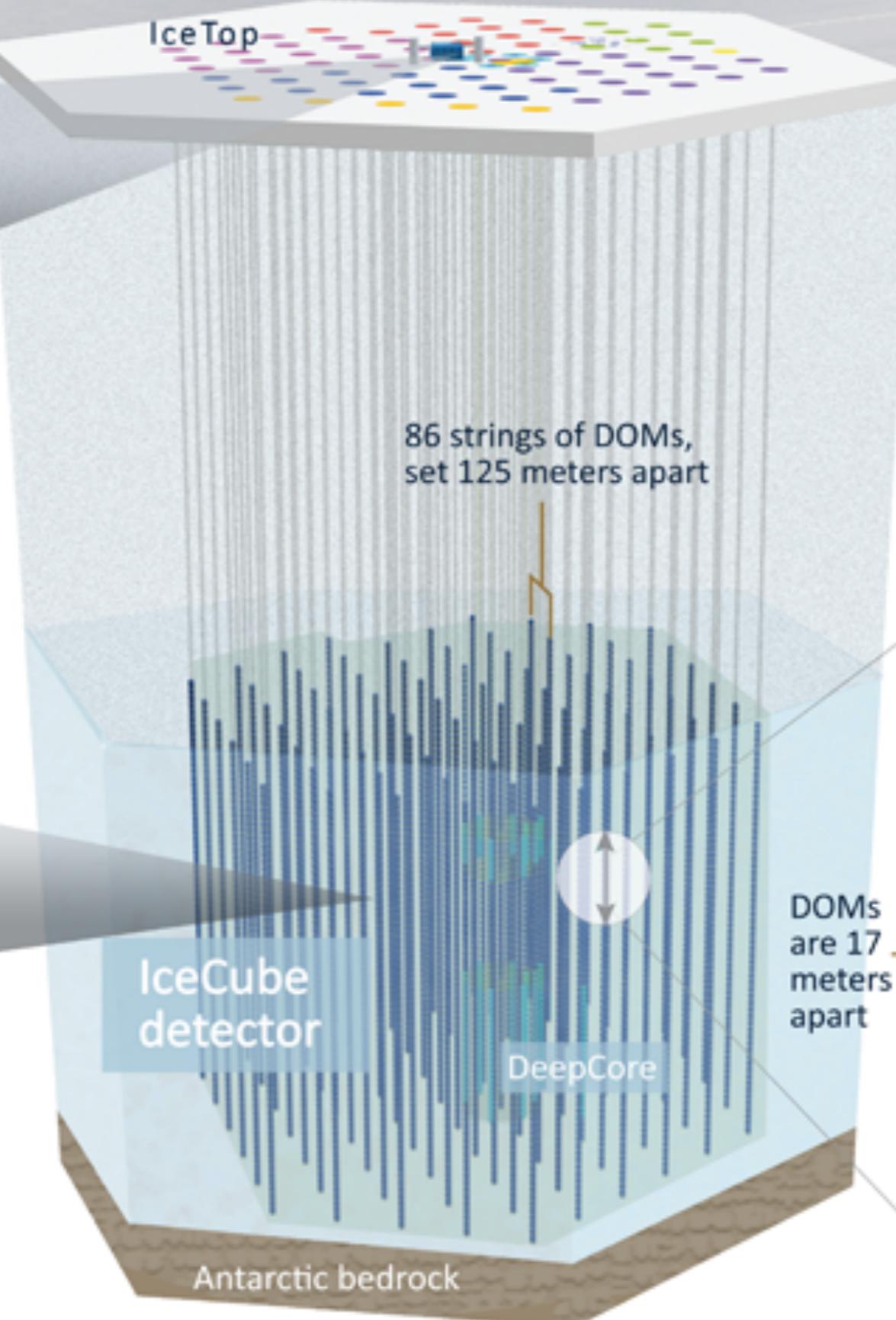
Digital Optical Module (DOM)

5,160 DOMs deployed in the ice

50 m

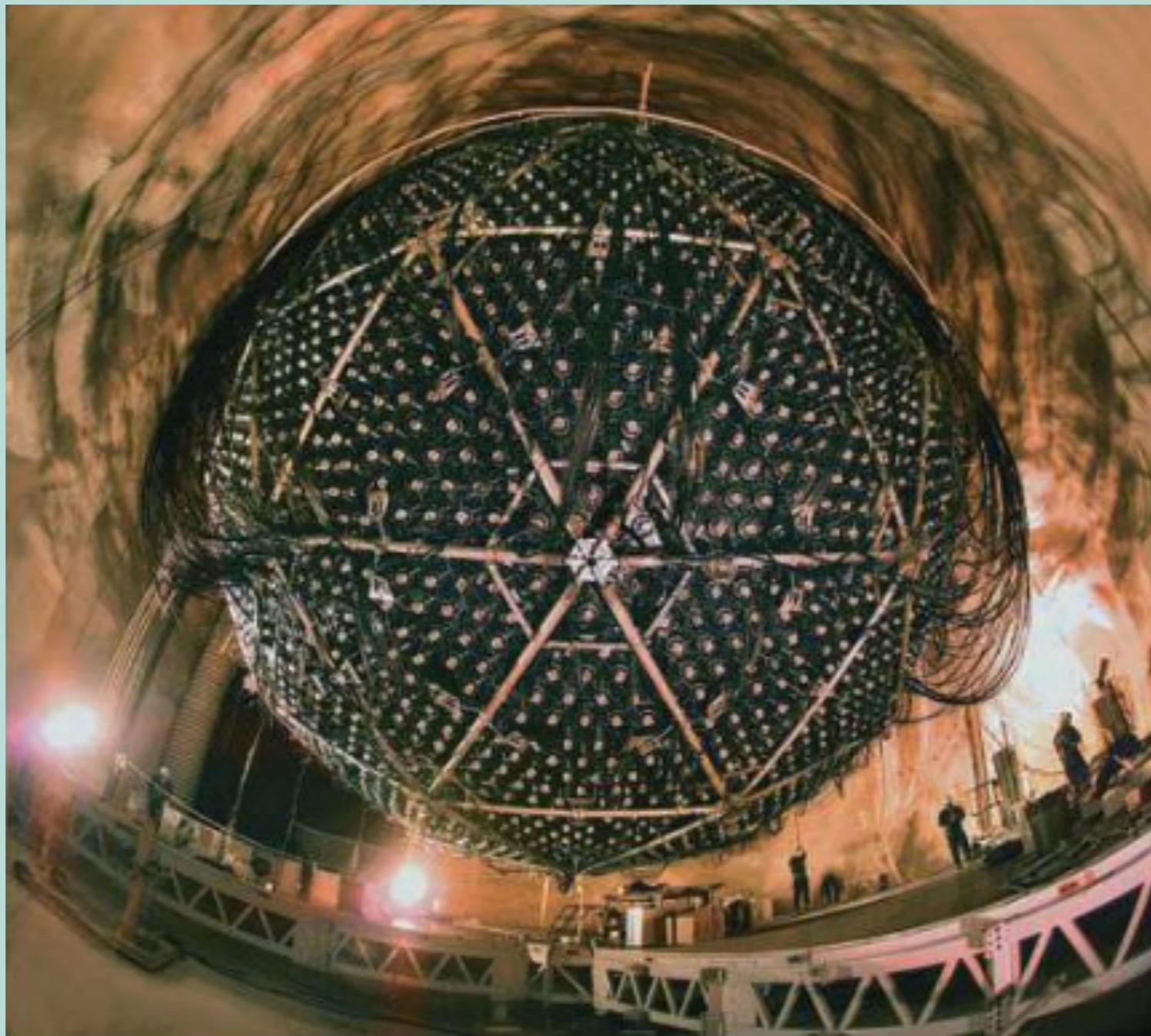
1450 m

2450 m



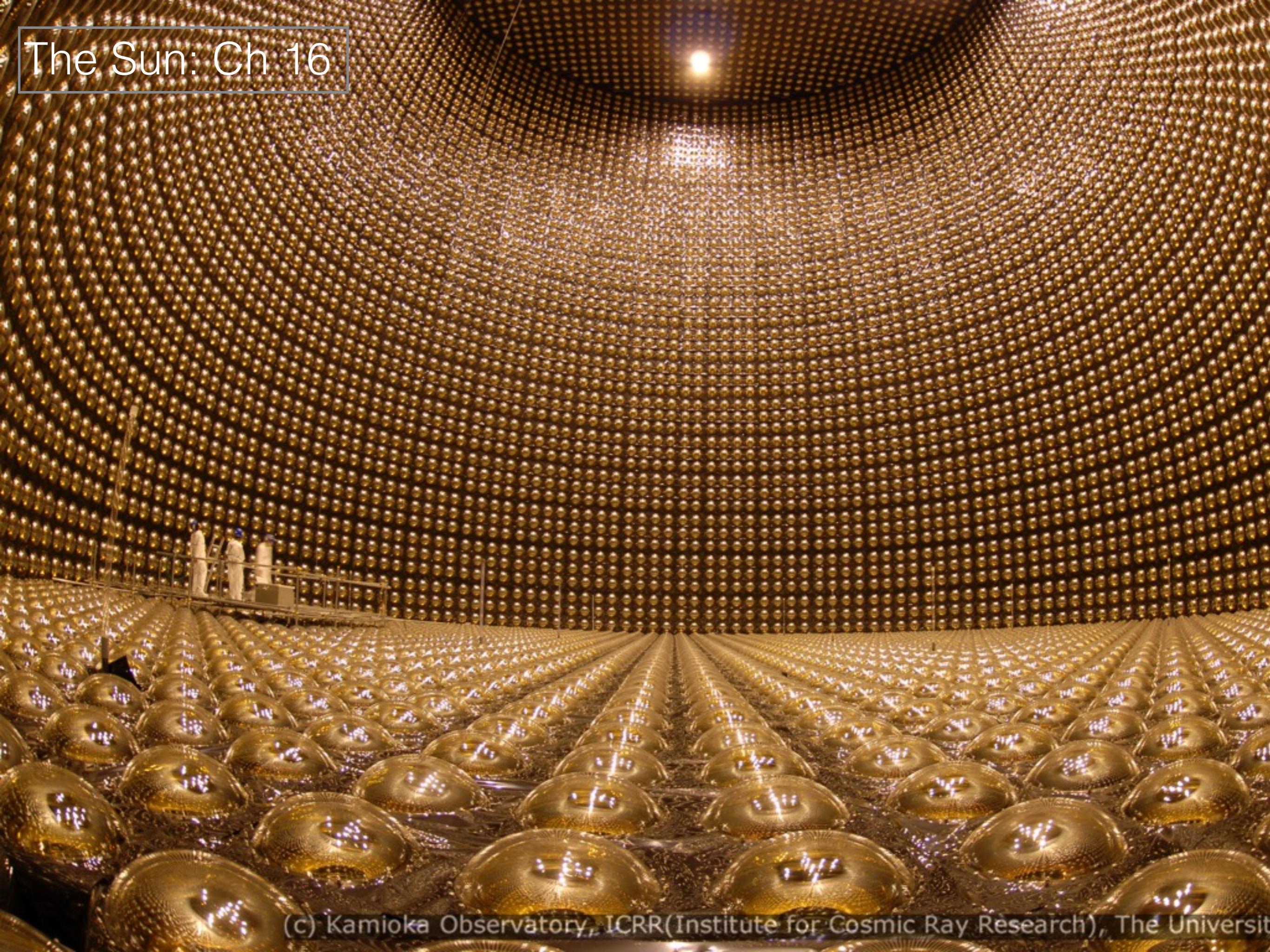
Amundsen–Scott South Pole Station, Antarctica
A National Science Foundation-managed research facility

The Sun: Ch 16



Sudbury Neutrino Solar Laboratory

The Sun: Ch 16



(c) Kamioka Observatory, ICRR (Institute for Cosmic Ray Research), The University of Tokyo