



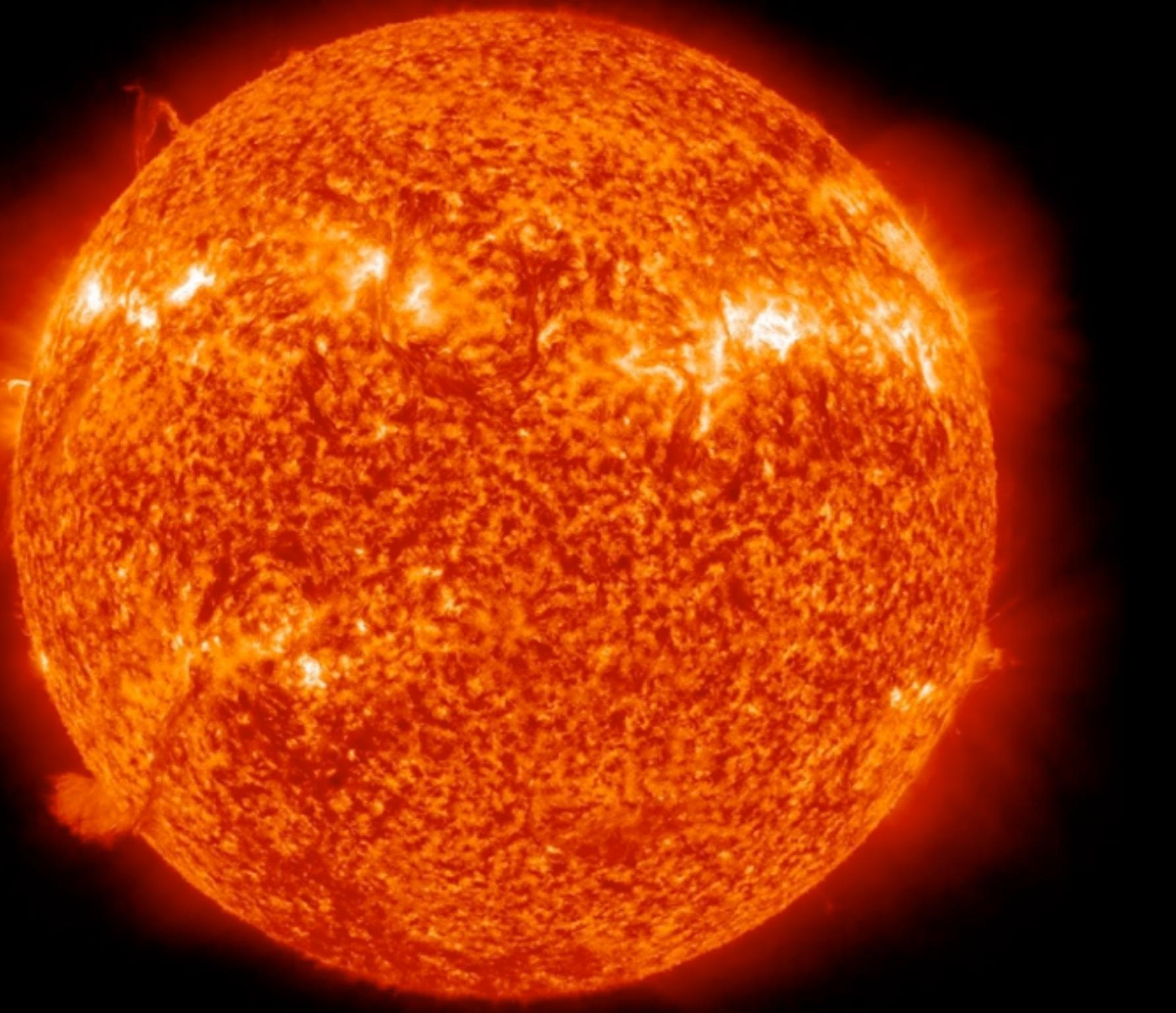
Lecture 14: The Sun



with your host:

Coop

The Sun



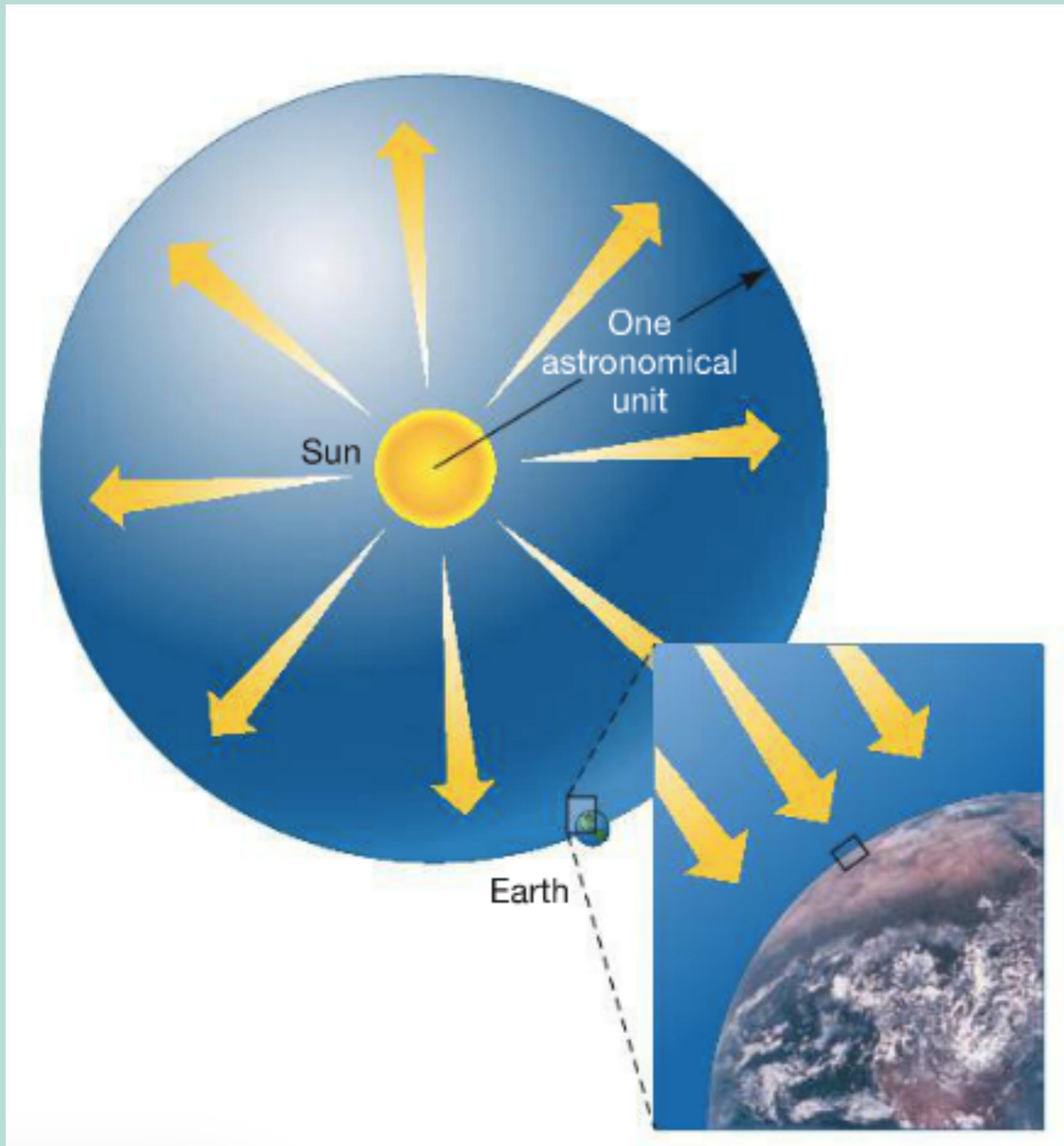
$m \sim 2\text{e}30 \text{ kg}$
 $\sim 300,000 m_e$

$r \sim 7\text{e}5 \text{ km}$
 $\sim 110 r_e$

$d \sim 28,000$
years

density ~
 1400 kg/m^3

- 0.5 degrees in the sky
- Differential Rotation
- 25 day rotation period at the equator, 36 at the poles
- Effective temperature of 5800 K



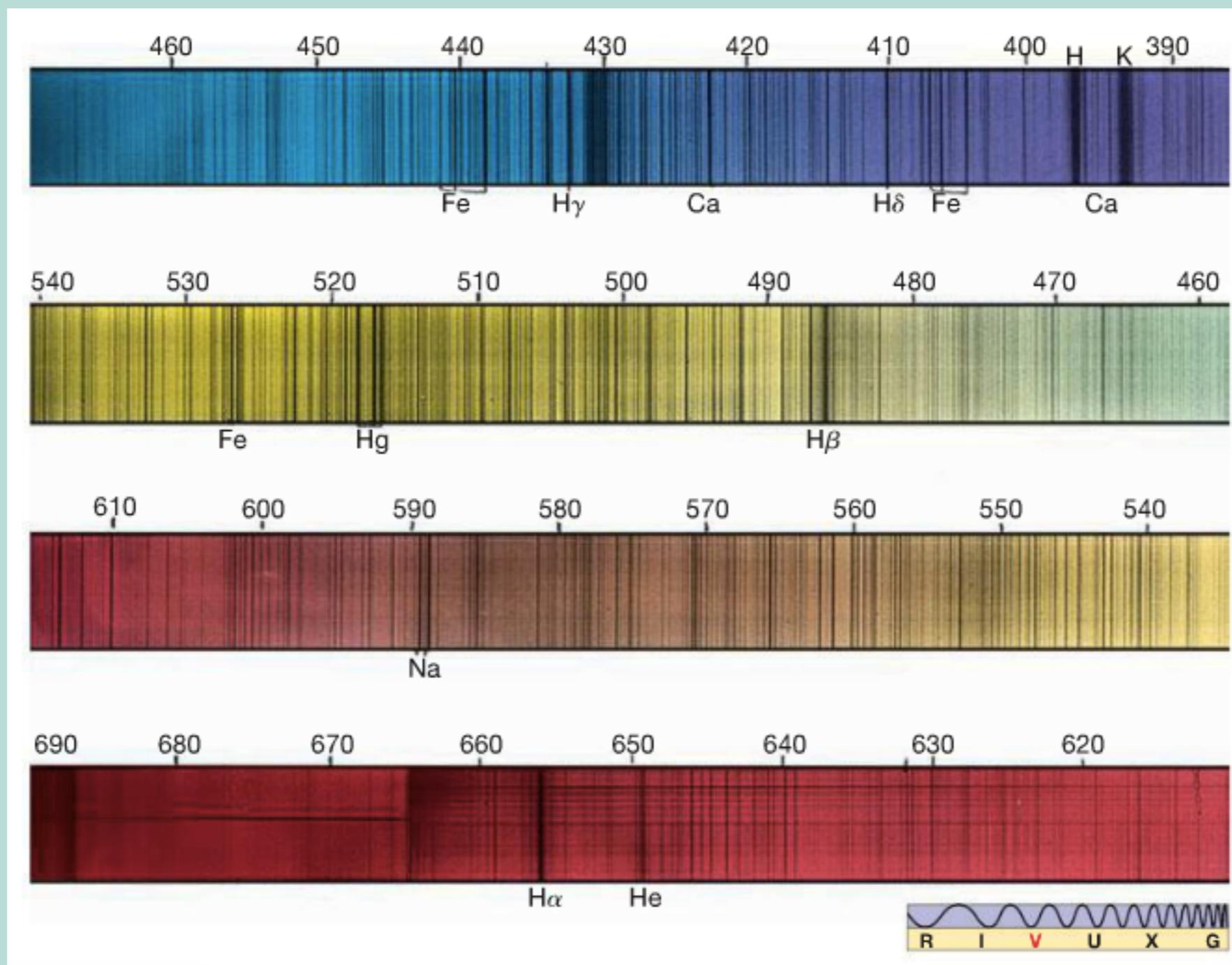
Just how bright is the sun?

- Solar constant: 1400 W/m^2
- 50% reaches earth, 30% goes to the atmosphere, 20% reflected away
 - About a mol of square meters in an AU sphere around the sun
 - Result: $4 \times 10^{26} \text{ W}$
 - 10 billion 1 megaton bombs/second
 - Could evaporate the world's oceans in 6 seconds, melt the crust in 3 minutes

TABLE 16.2 The Composition of the Sun

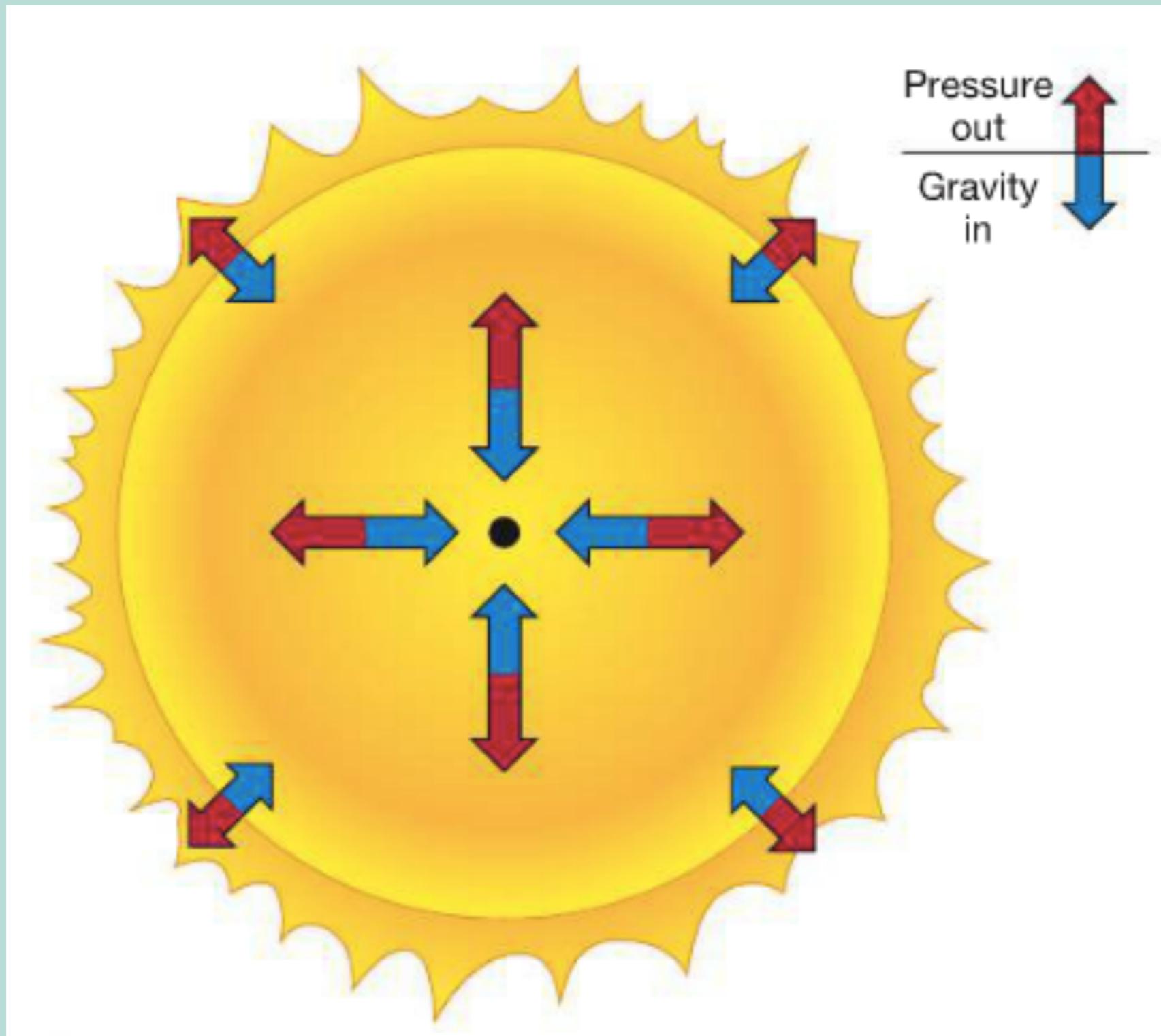
Element	Percentage of Total Number of Atoms	Percentage of Total Mass
Hydrogen	91.2	71.0
Helium	8.7	27.1
Oxygen	0.078	0.97
Carbon	0.043	0.40
Nitrogen	0.0088	0.096
Silicon	0.0045	0.099
Magnesium	0.0038	0.076
Neon	0.0035	0.058
Iron	0.0030	0.14
Sulfur	0.0015	0.040

Solar Spectrum

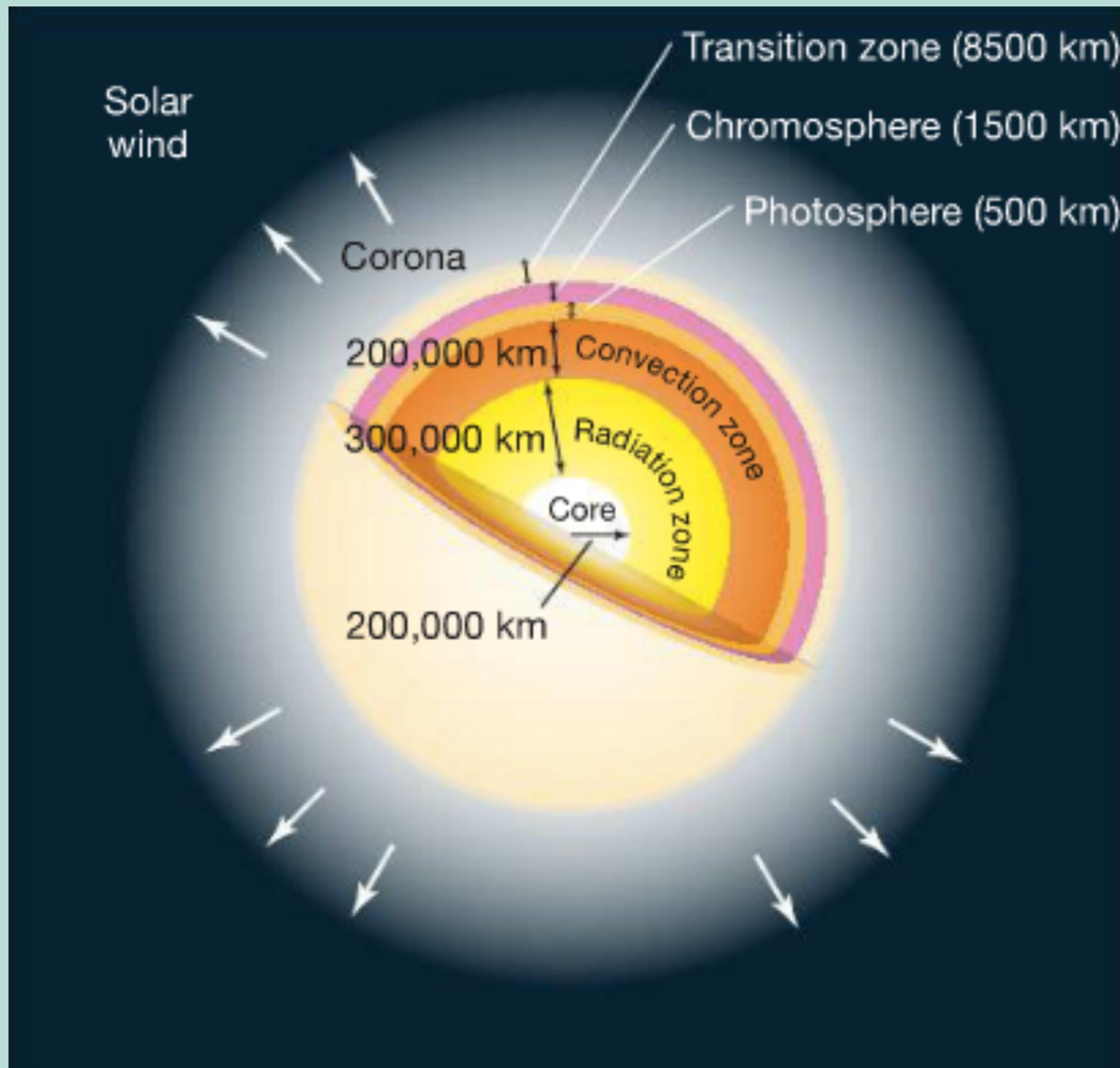


- Recall, we learn not just about the emitted radiation (emission lines), but also the intermediate atmosphere of the sun (absorption lines)

Hydrostatic Equilibrium



Standard Solar Model

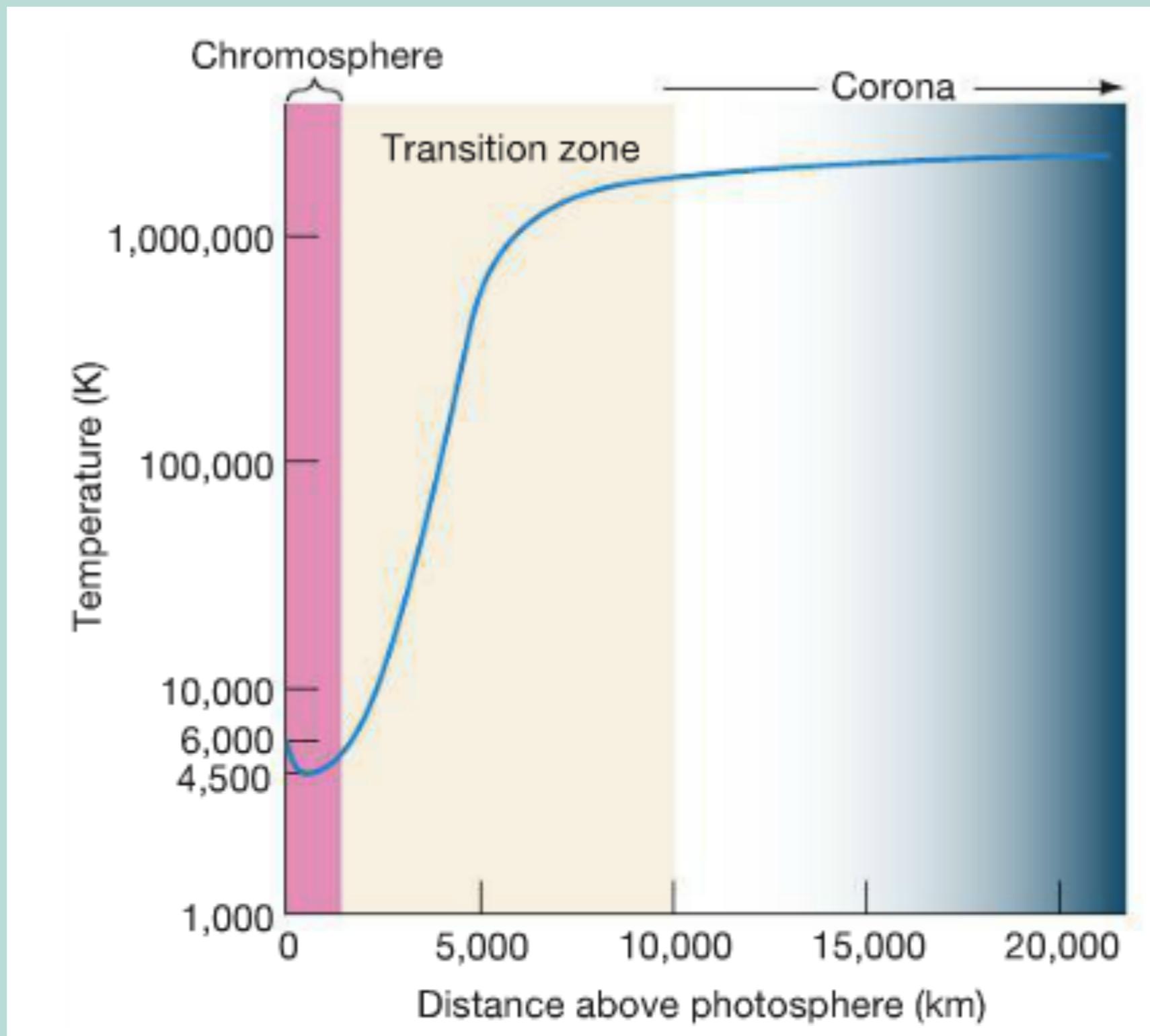


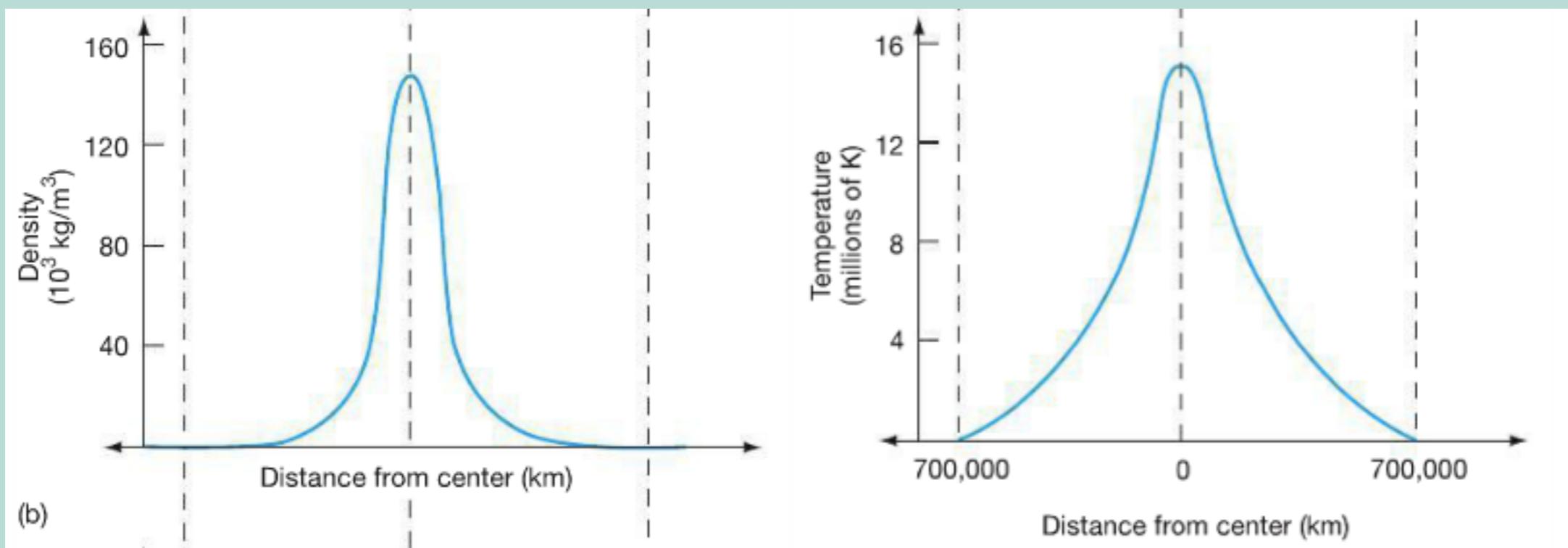
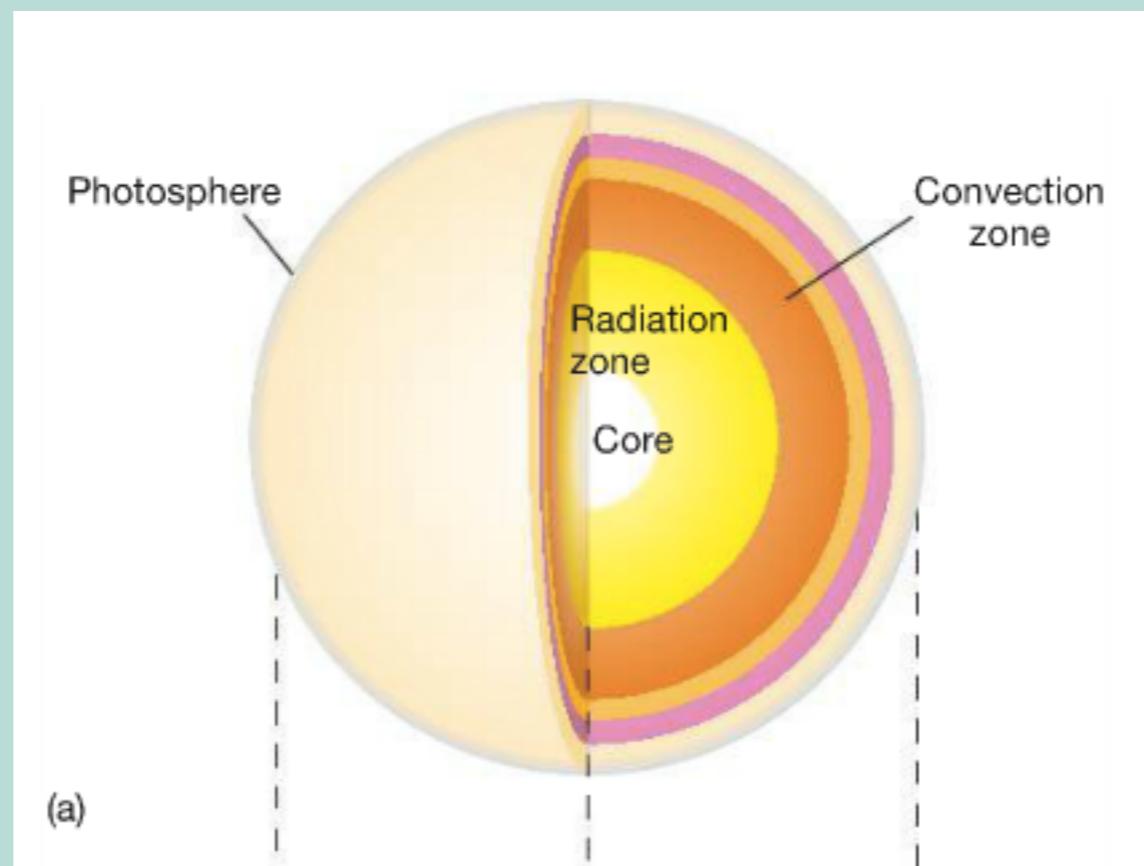
- In the core it's too hot to form atoms, so radiation doesn't get absorbed by anything (radiation zone)
- Further out, it's cool enough to form atoms (how cool?) and so ALL the light gets absorbed and turned into heat and reflected light
- This heat CONVECTS to the surface which glows like a black body (photosphere)

TABLE 16.1 The Standard Solar Model

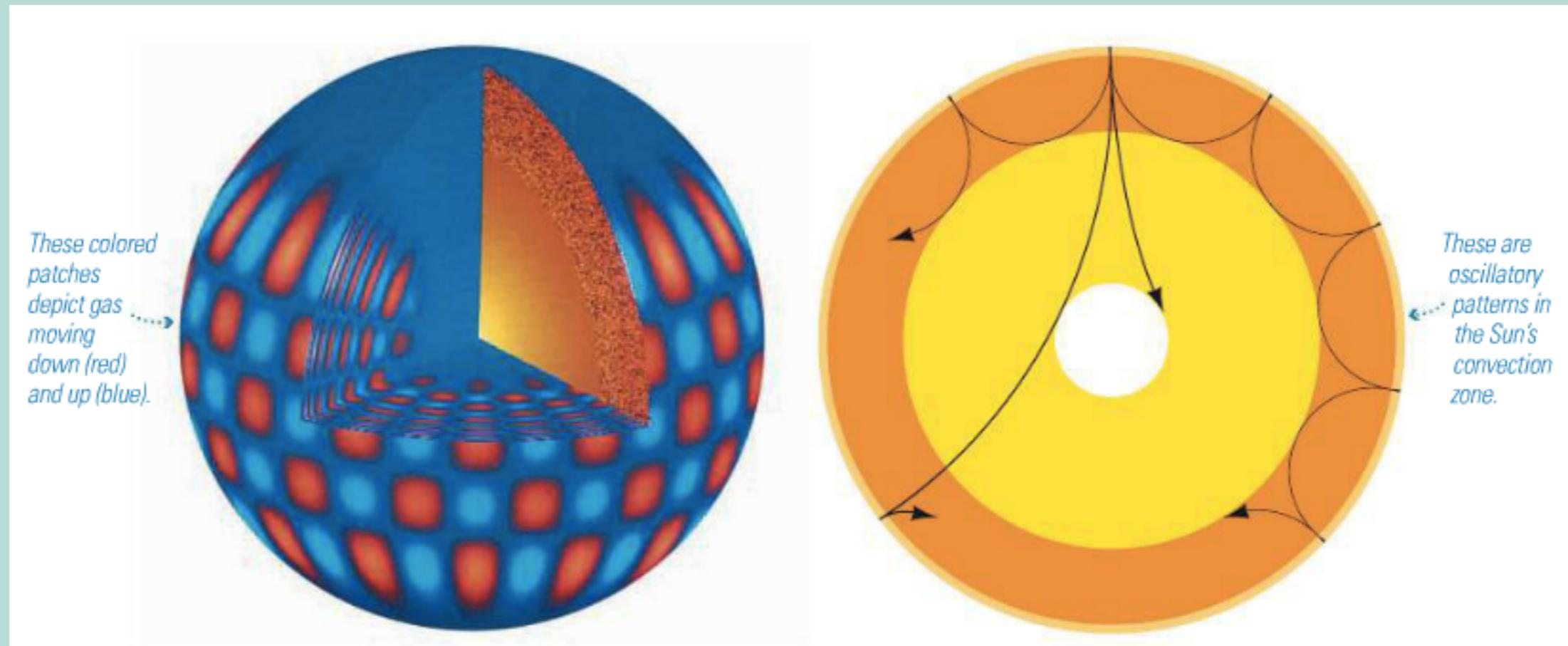
Region	Inner Radius (km)	Temperature (K)	Density (kg/m ³)	Defining Properties
Core	0	15,000,000	150,000	Energy generated by nuclear fusion
Radiation zone	200,000	7,000,000	15,000	Energy transported by electromagnetic radiation
Convection zone	496,000*	2,000,000	150	Energy carried by convection
Photosphere	696,000*	5800	2×10^{-4}	Electromagnetic radiation can escape—the part of the Sun we see
Chromosphere	696,500*	4500	5×10^{-6}	Cool lower atmosphere
Transition zone	698,000*	8000	2×10^{-10}	Rapid increase in temperature
Corona	706,000*	3,000,000	10^{-12}	Hot, low-density upper atmosphere
Solar wind	10,000,000	> 1,000,000	10^{-23}	Solar material escapes into space and flows outward through the solar system

Atmospheric Physics is Hard



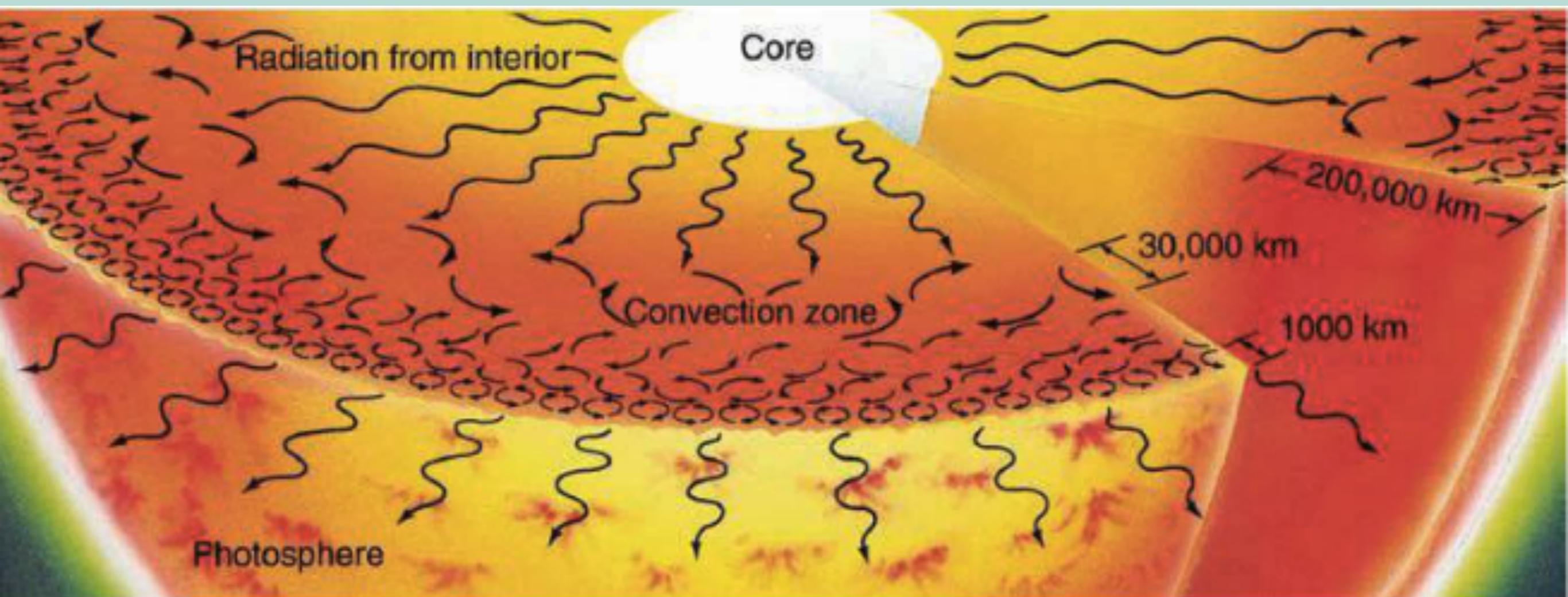


Helioseismology



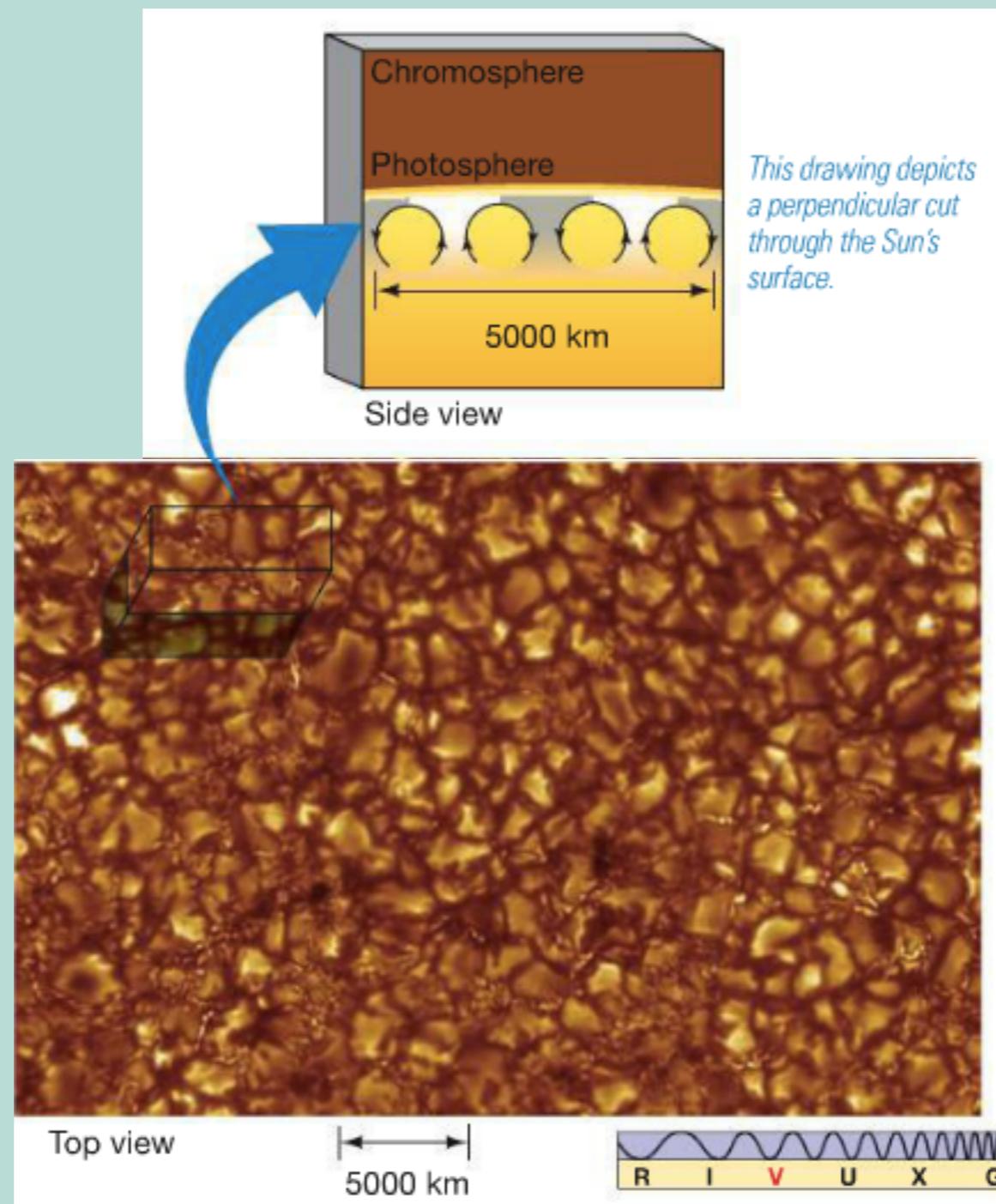
Global Oscillations Network Group (GONG)

Convection Cells



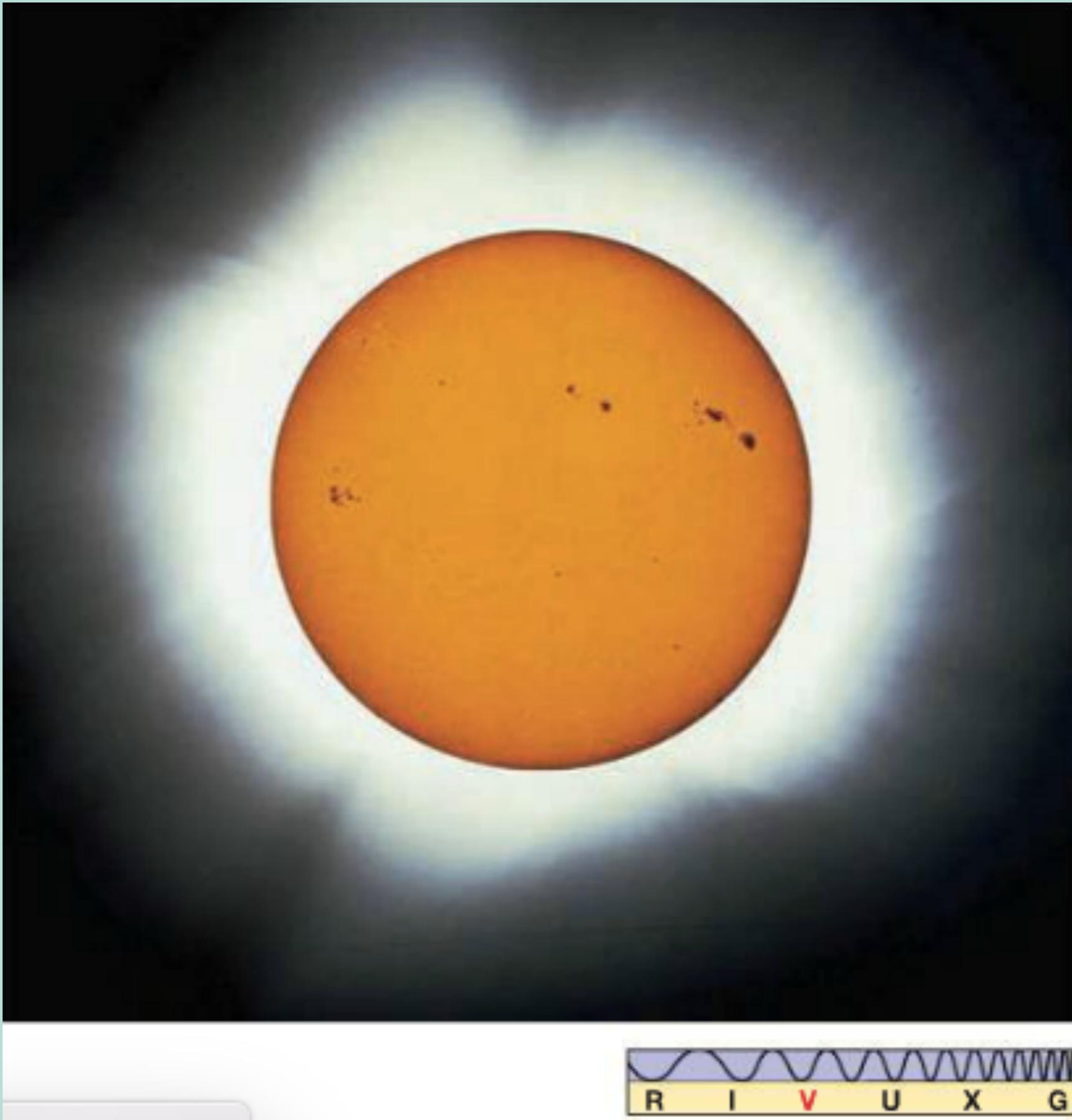
Hierarchy of cells

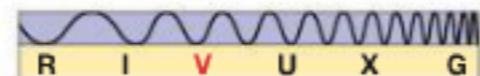
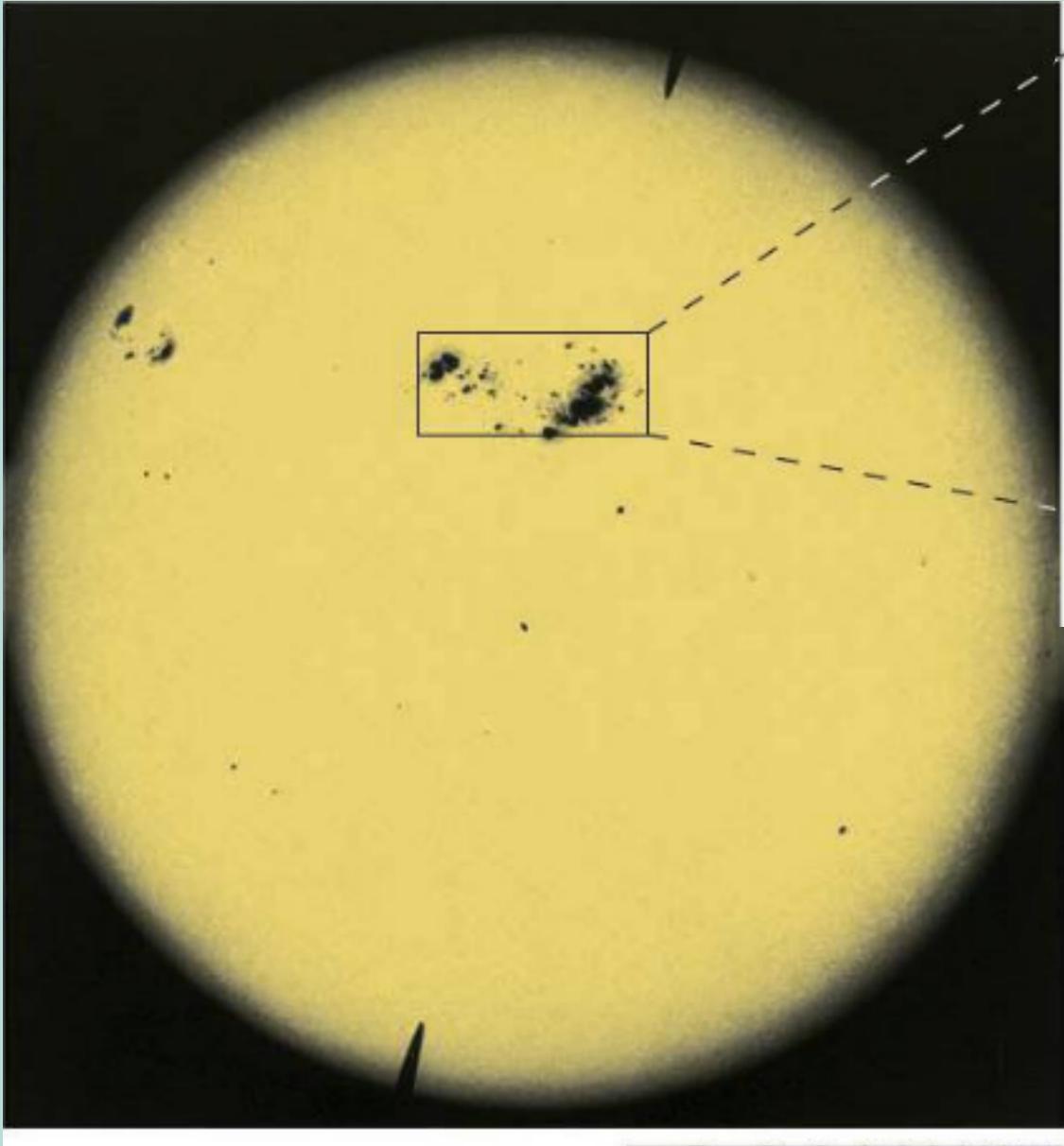
Solar Granules



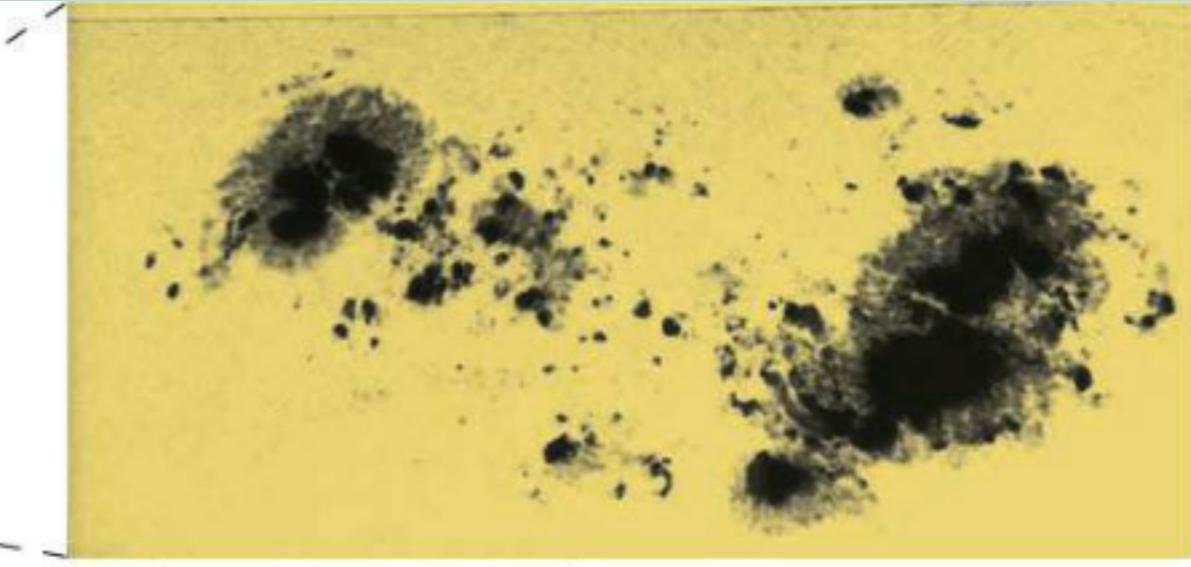
- Each granule is approximately the size of the Earth!

Sun Spots from Earth

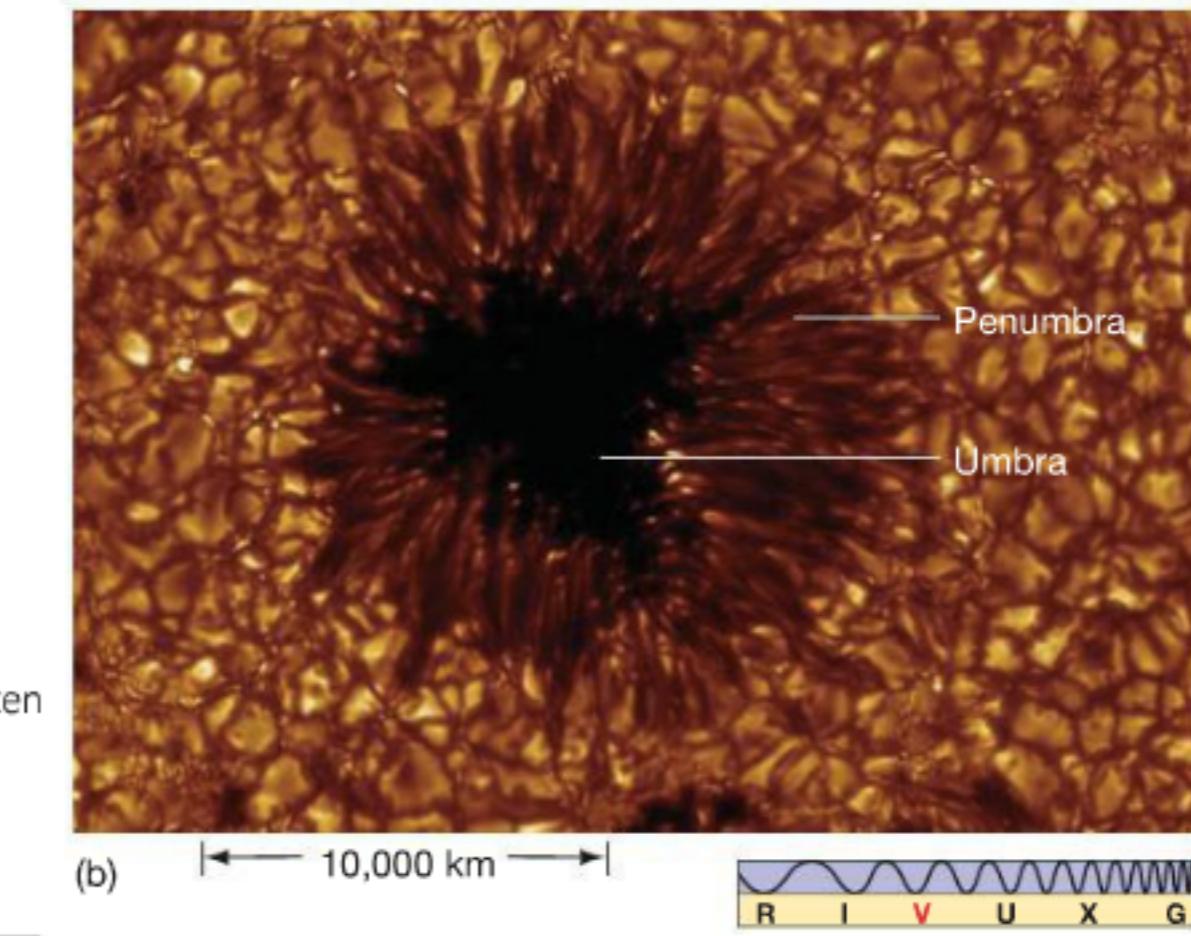




Sun Spots

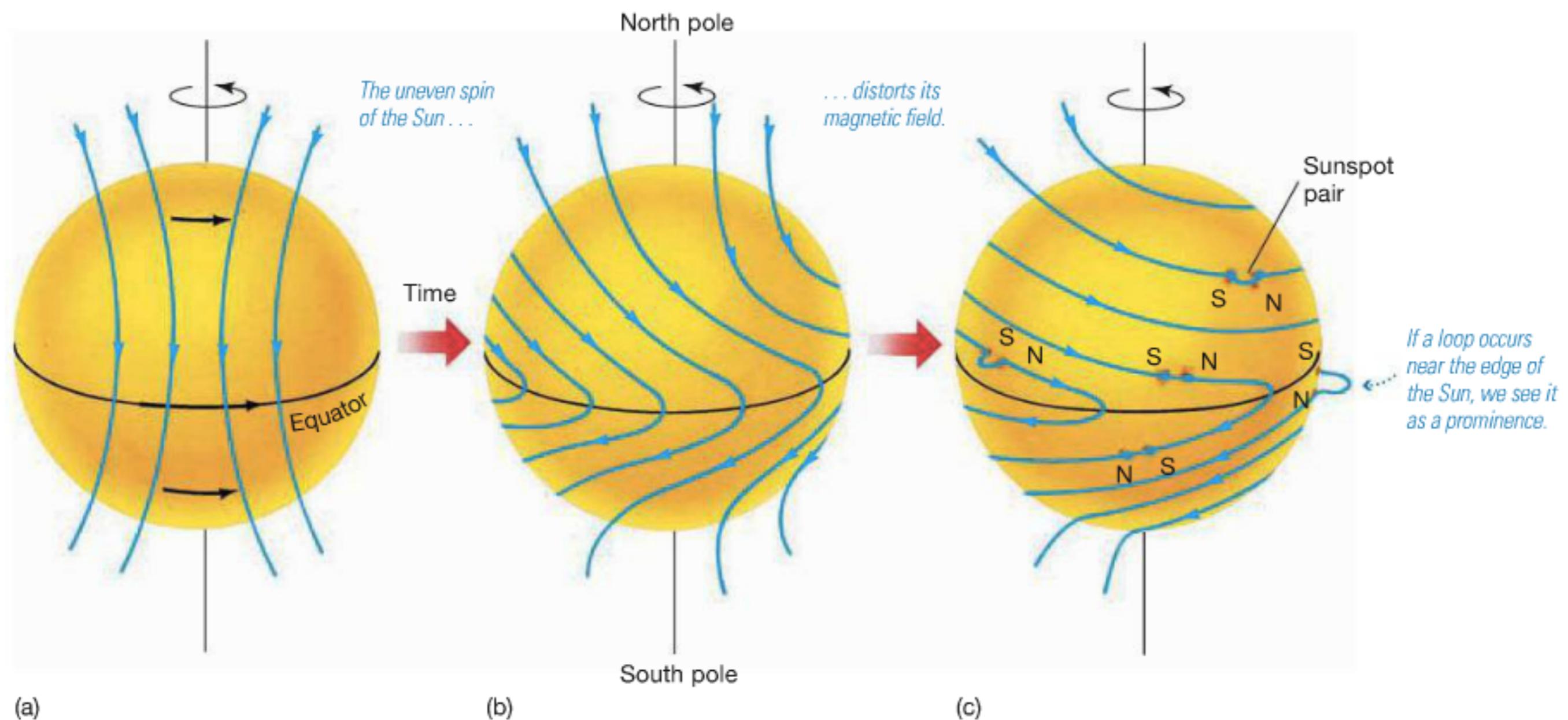


Sunspots appear dark because they are slightly cooler than the surrounding gas.

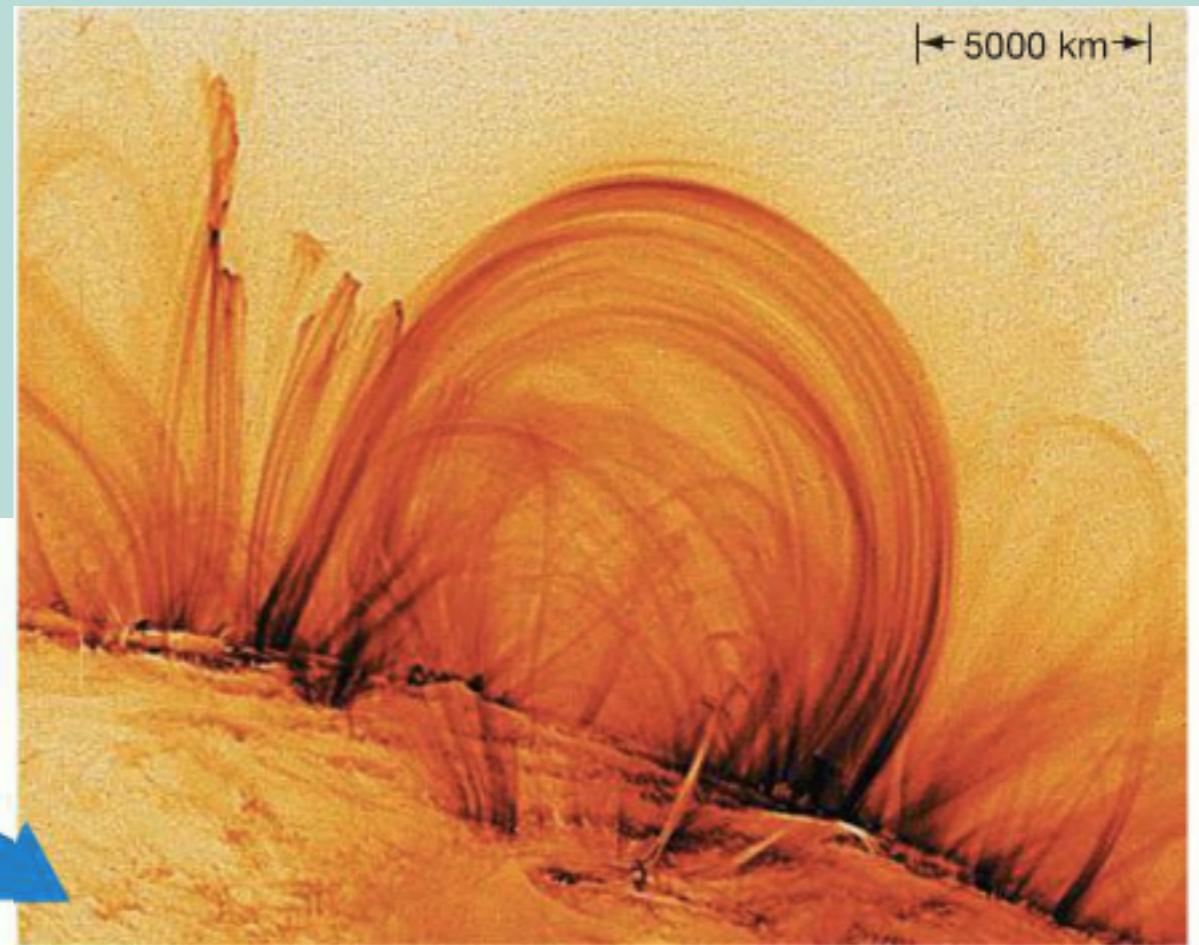
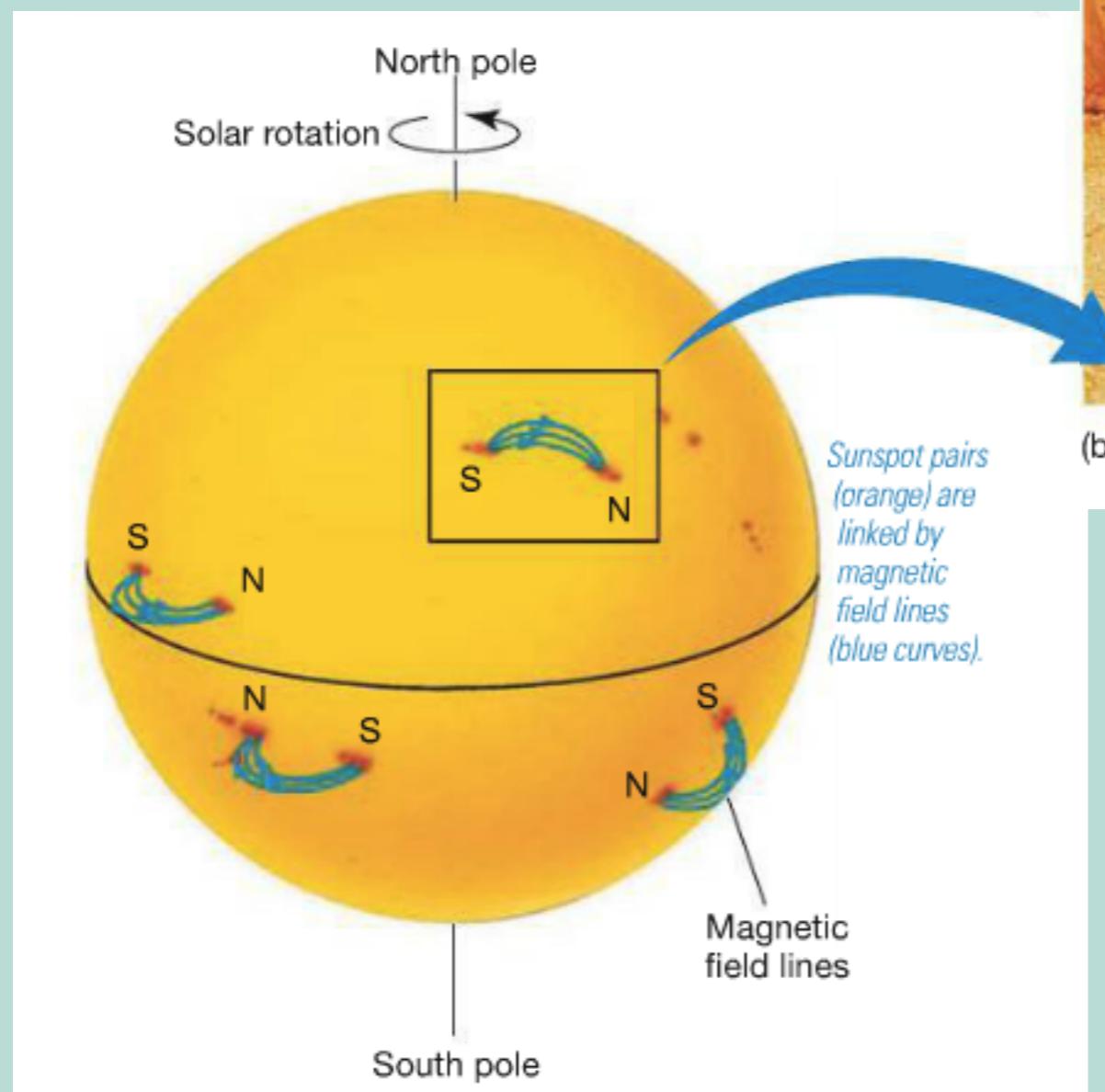


- Umbra is 1000K cooler than the rest of the sun (still bright!)

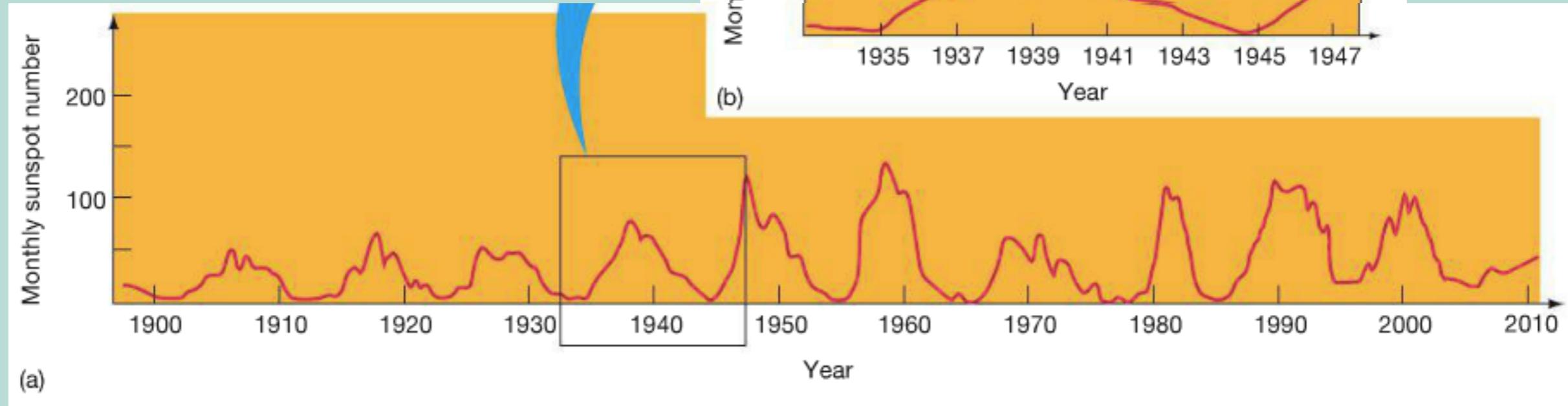
Origin of Sun Spots

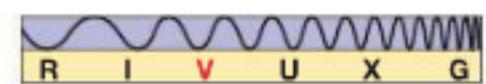
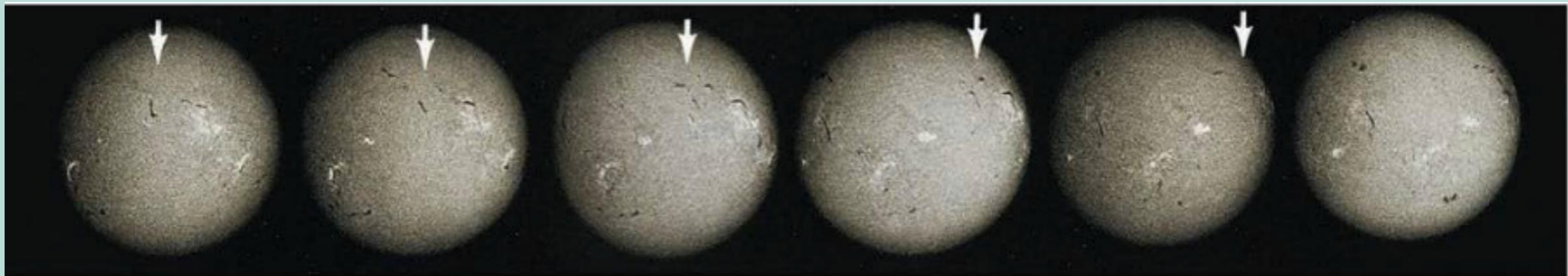
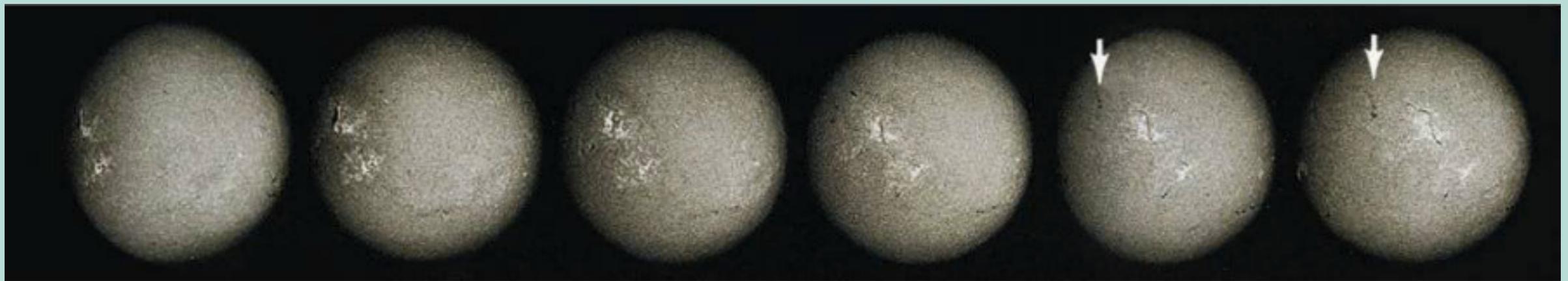


Sun Spots and Magnetism



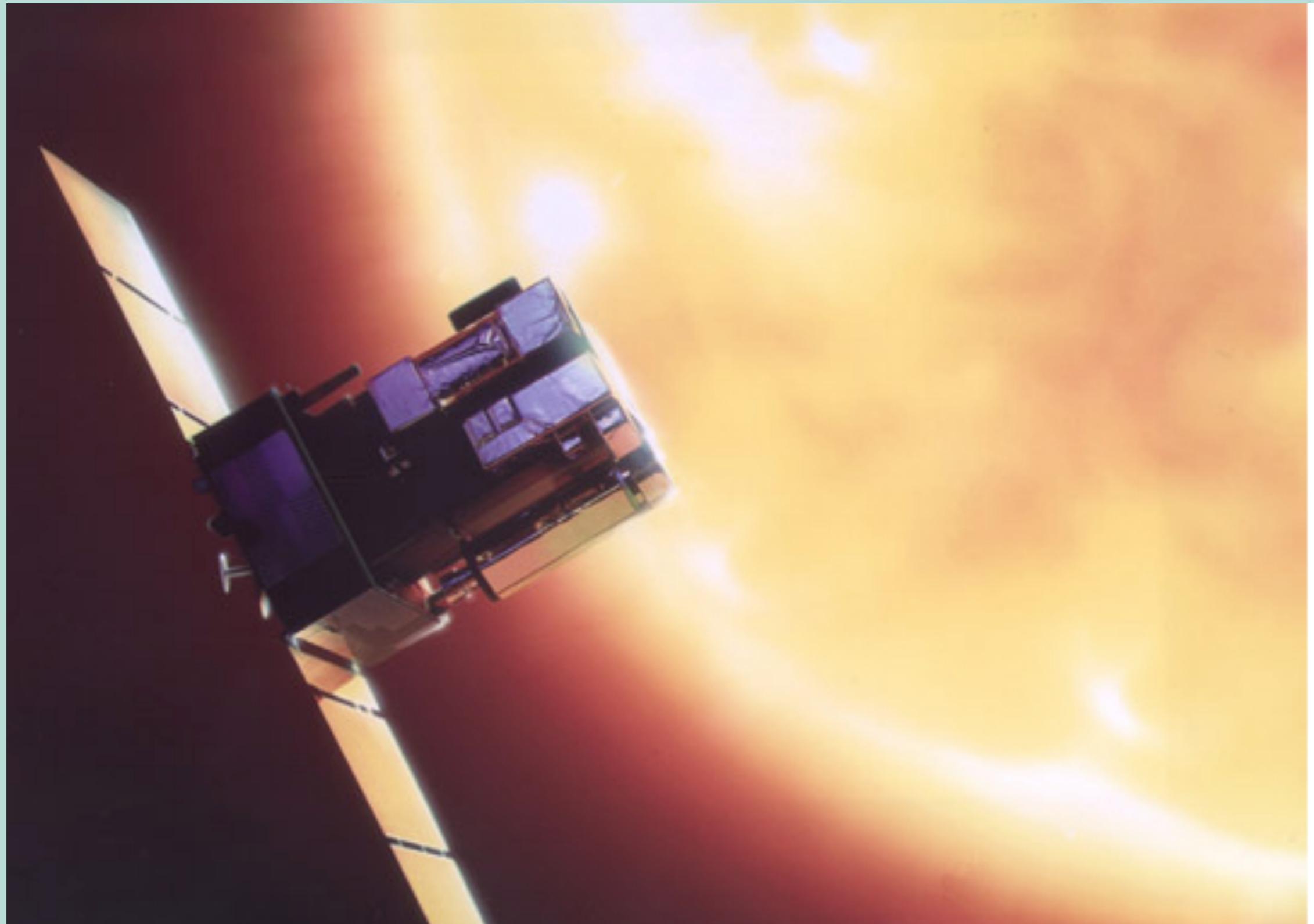
Solar Cycle: 11 + 11 years





Sun Spots in Motion

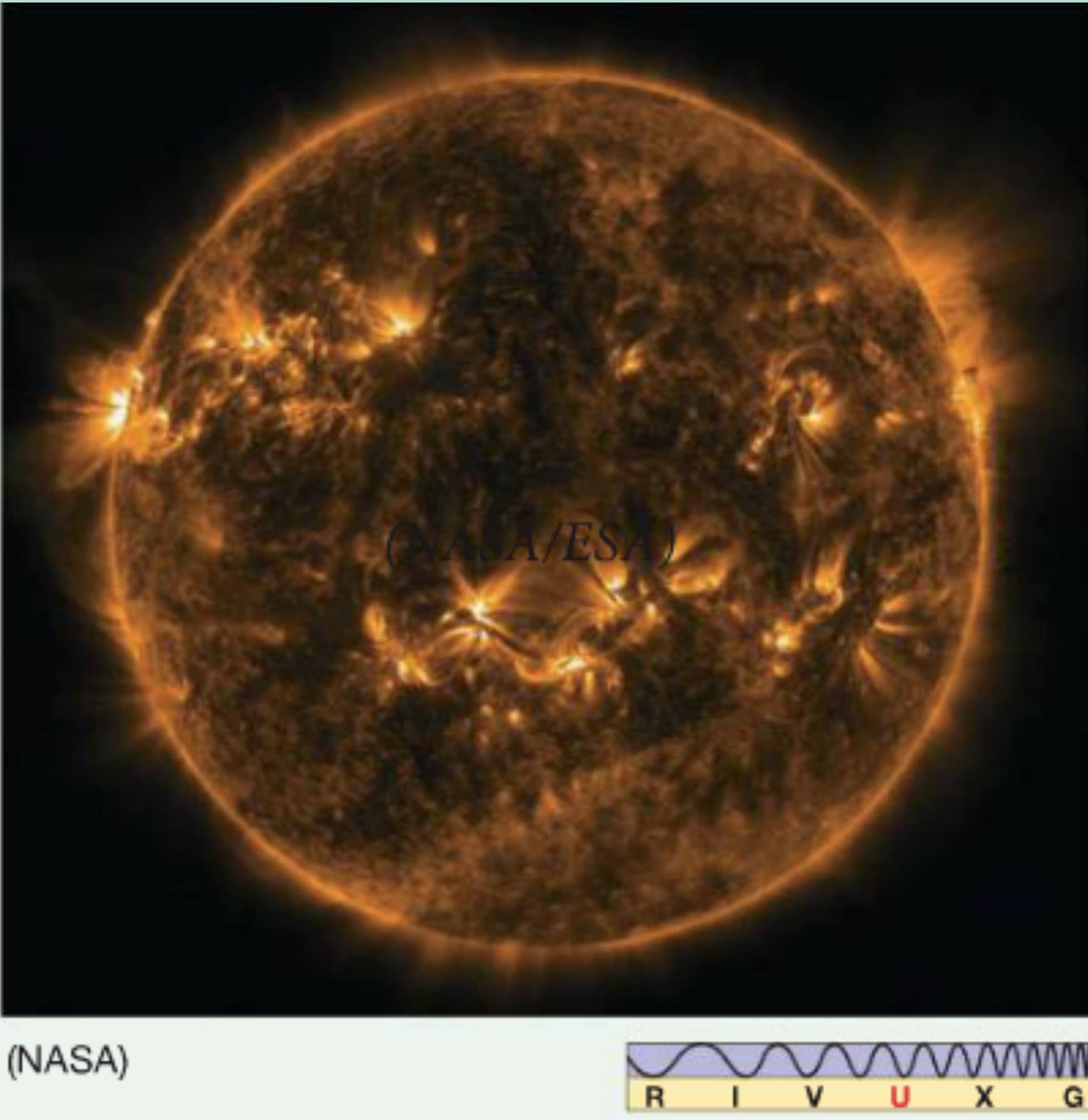
SOHO (SOlar Heliospheric Observatory)



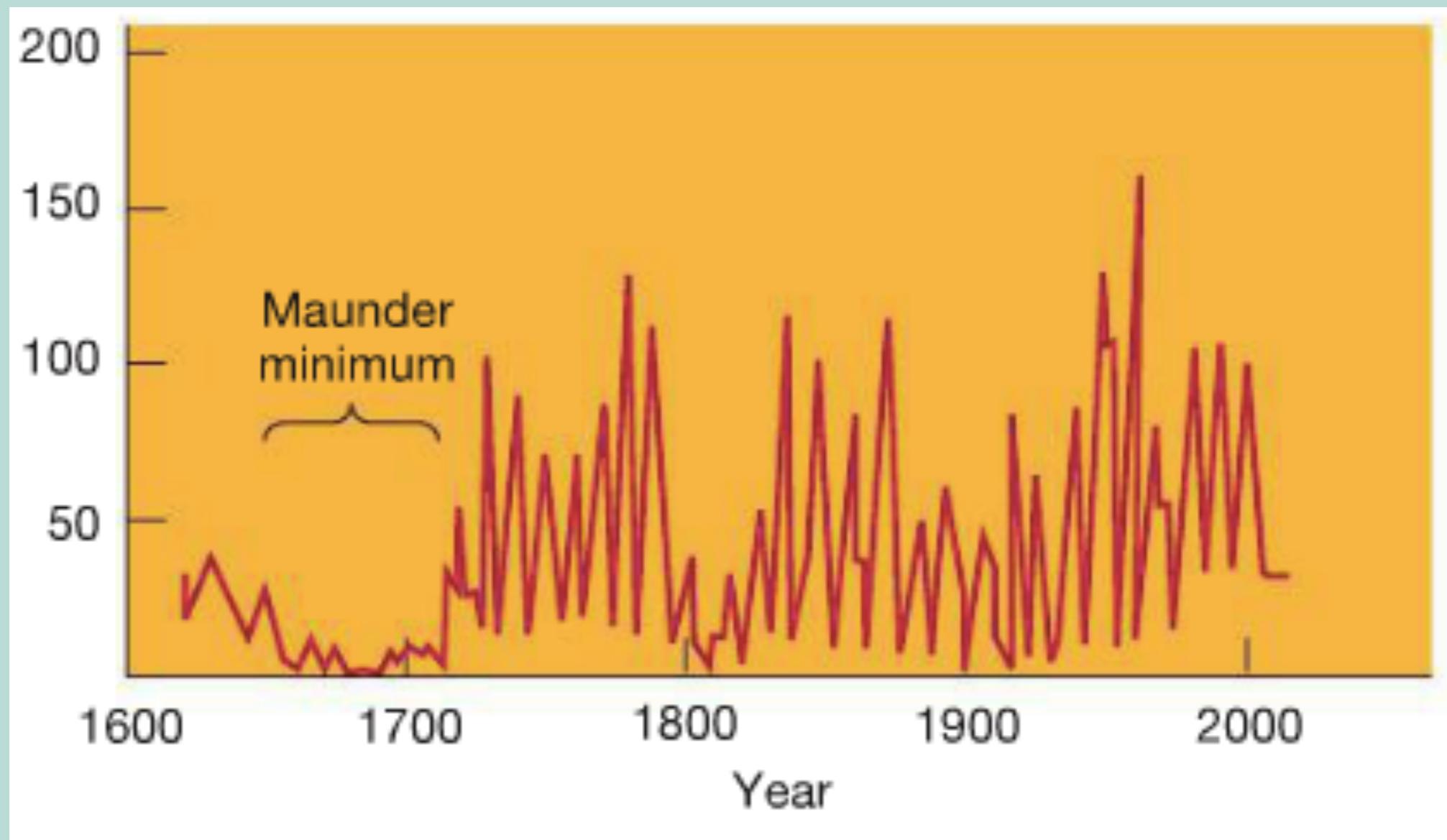
(L1 point)

LINK

SDO (Solar Dynamics Observatory)



Maunder Minimum



A beautiful summer day in the Netherlands...

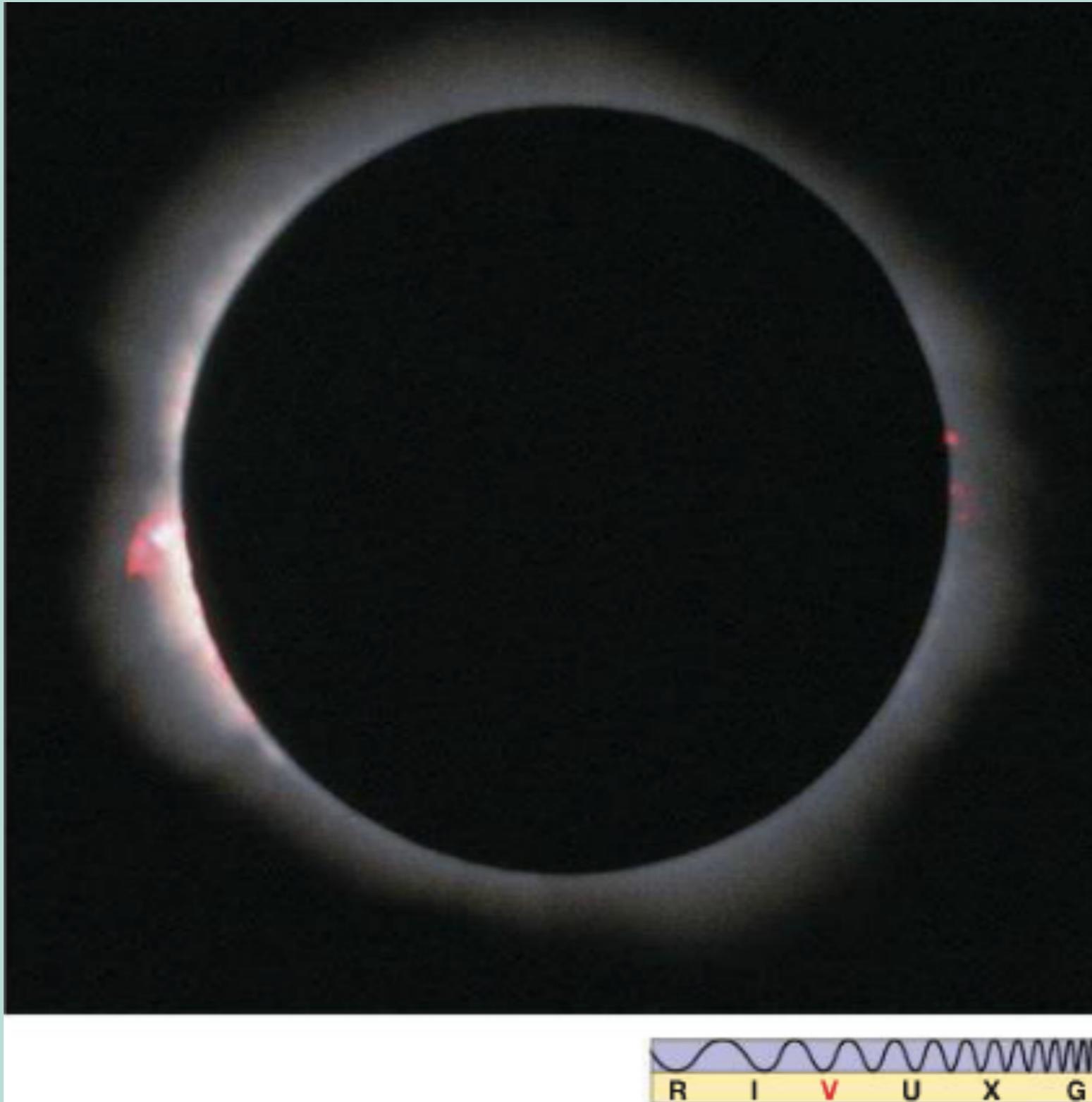


No real evidence of weather



Solar Activity

Chromosphere



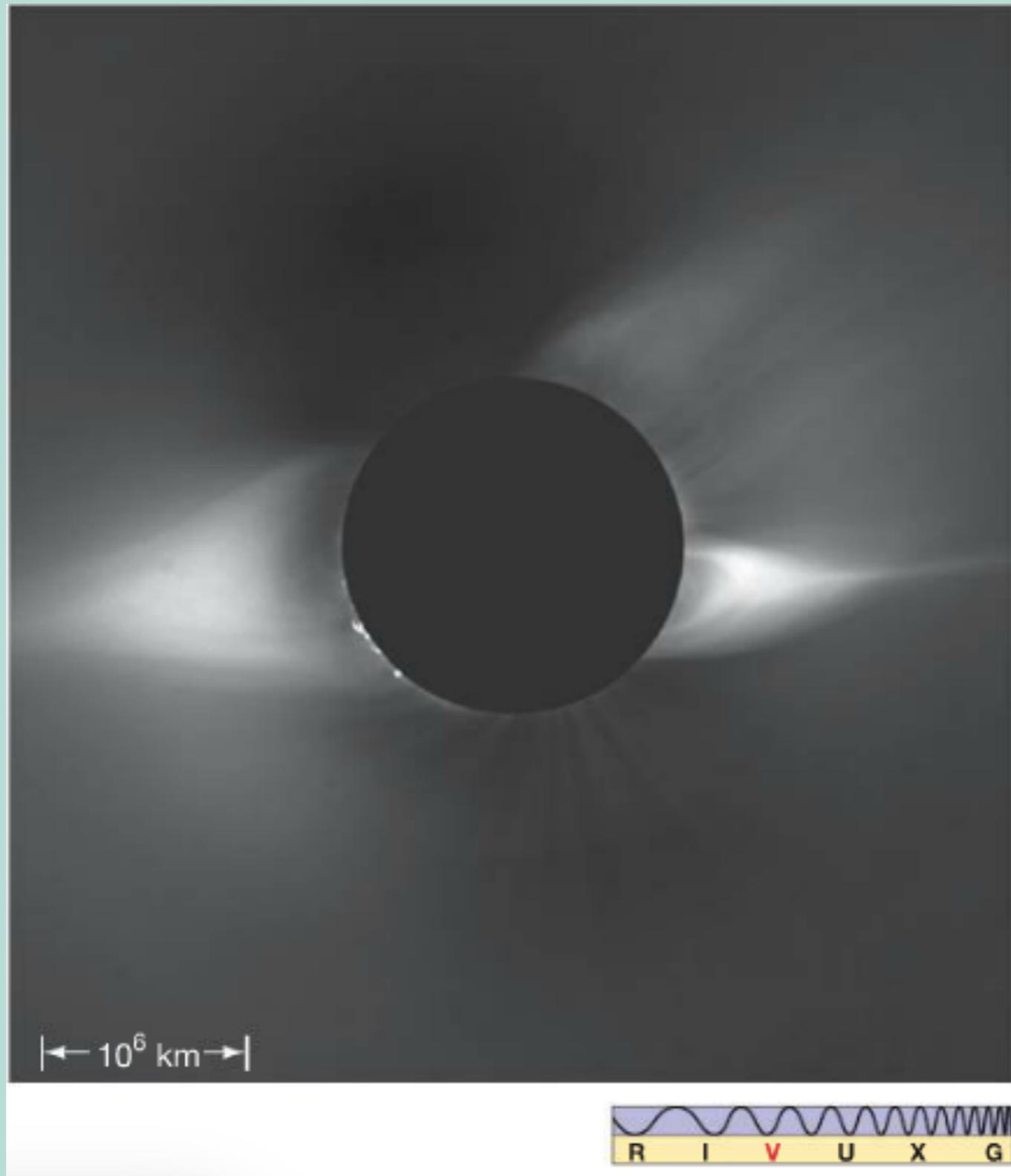
(red hue due to Ha lines)

Corona

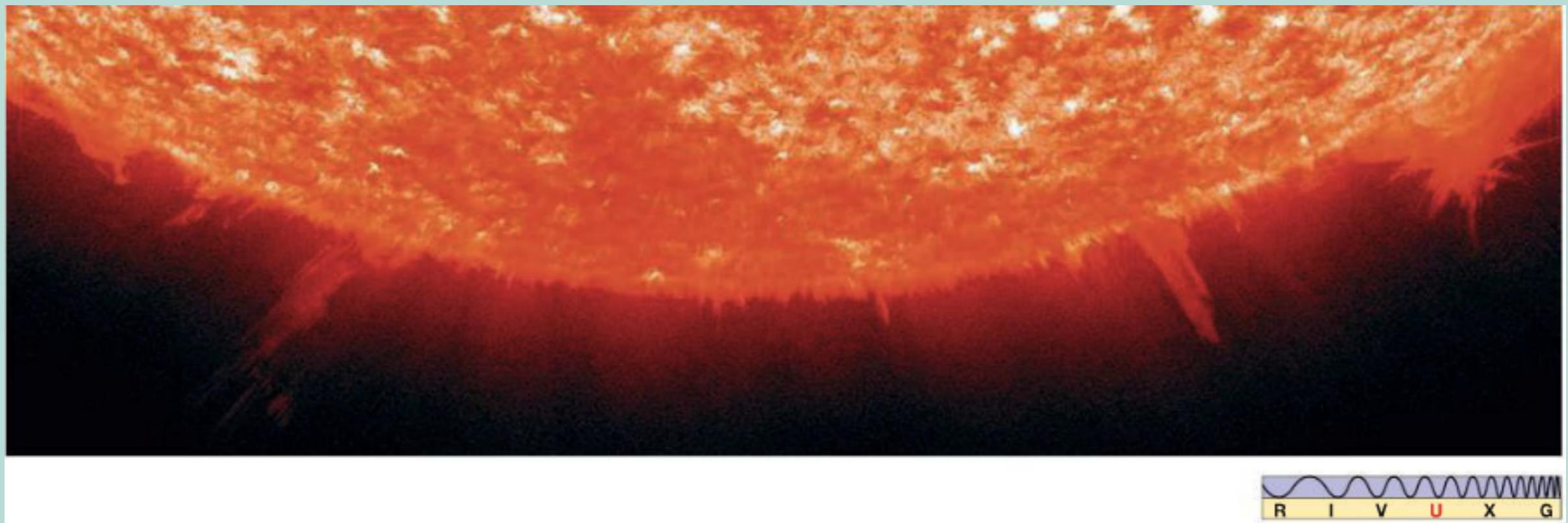


- Solar wind blasts out towards earth at 500 km/s
- Thus the sun is “evaporating” at a rate of 2 million tons of solar matter per second
- Sounds big, but amounts to 0.1% of its mass over the last 4.6 billion years

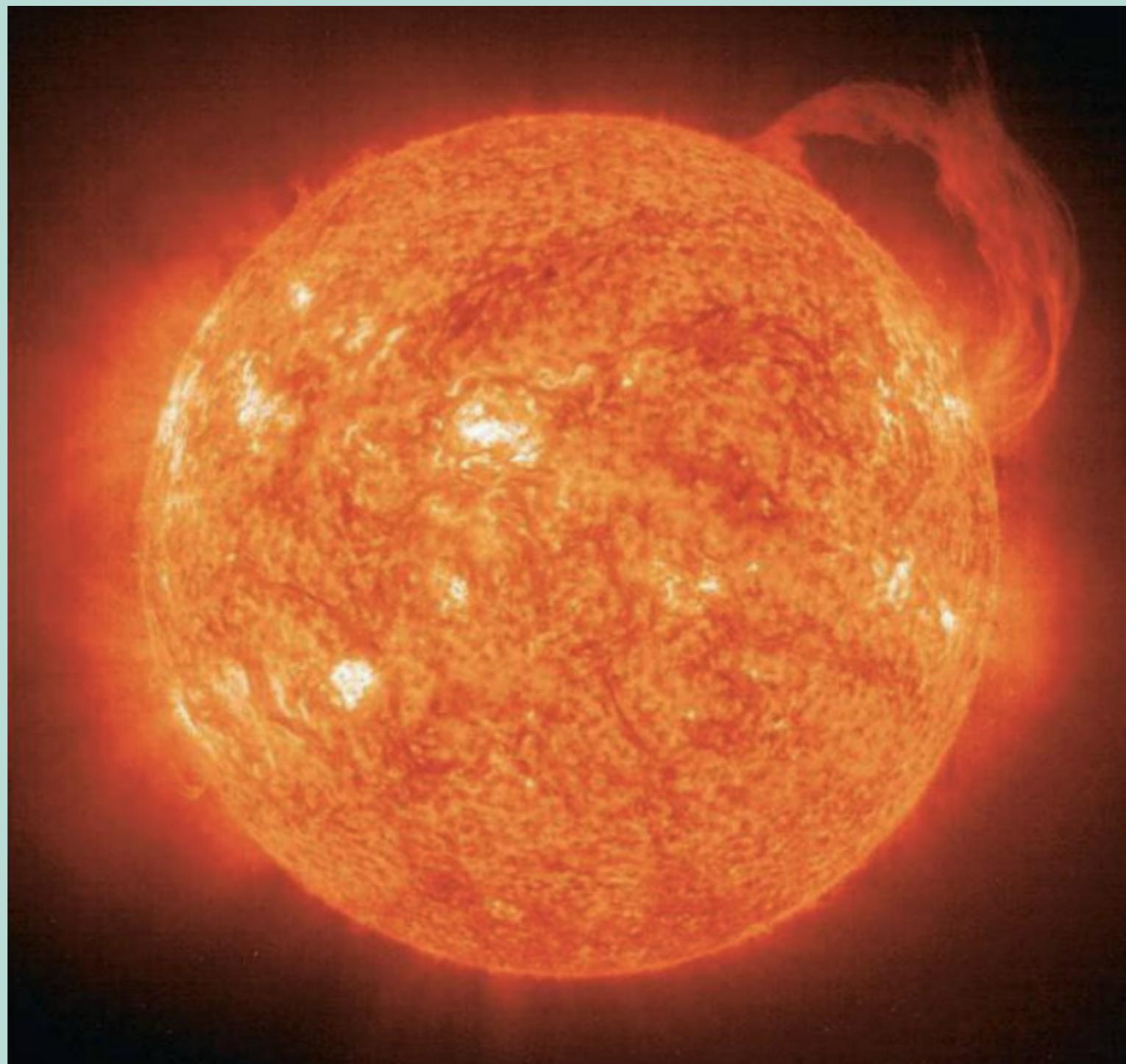
Corona during solar cycle peak



Solar Spicules



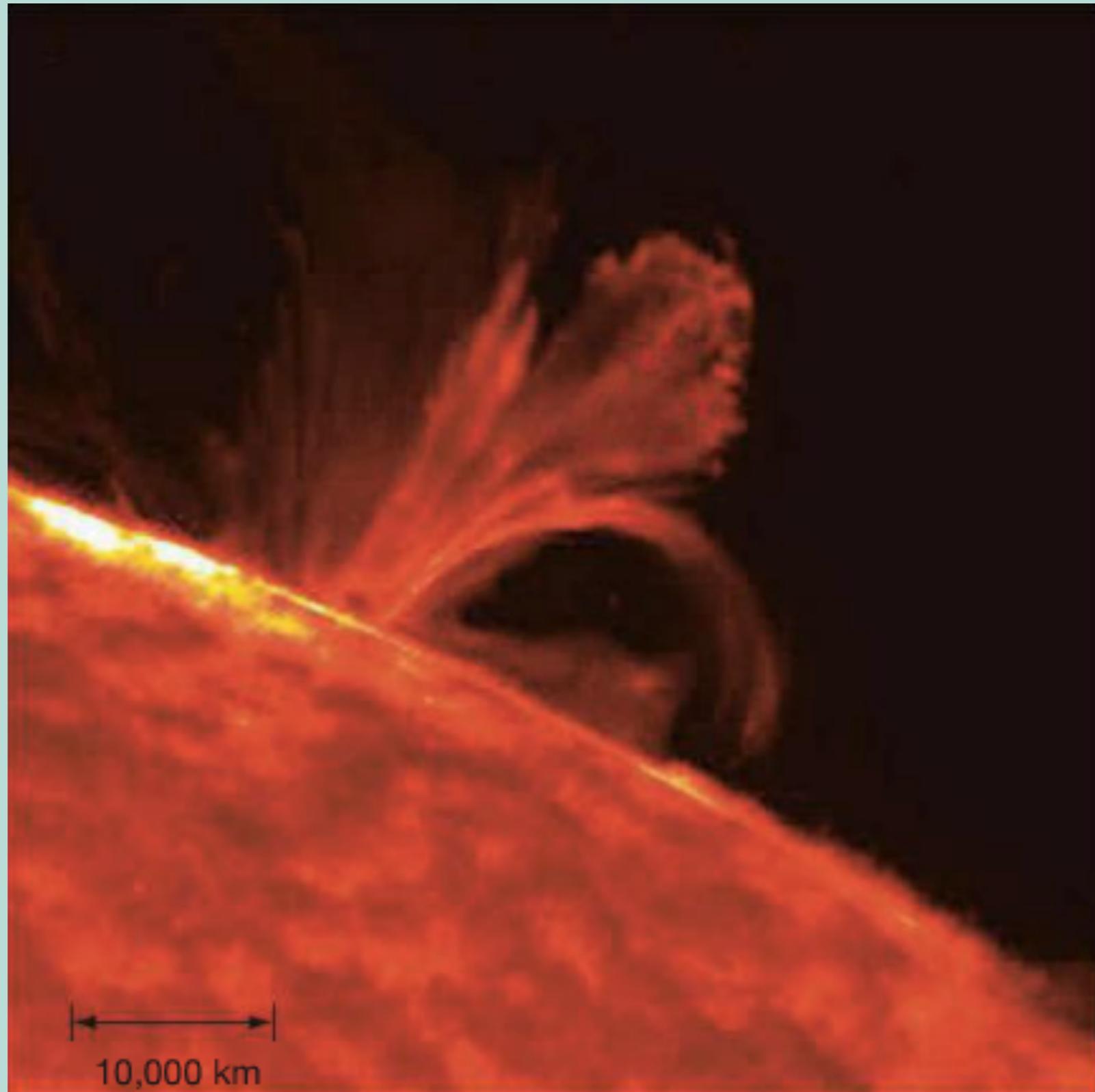
Serious Solar Prominence



Solar Prominence



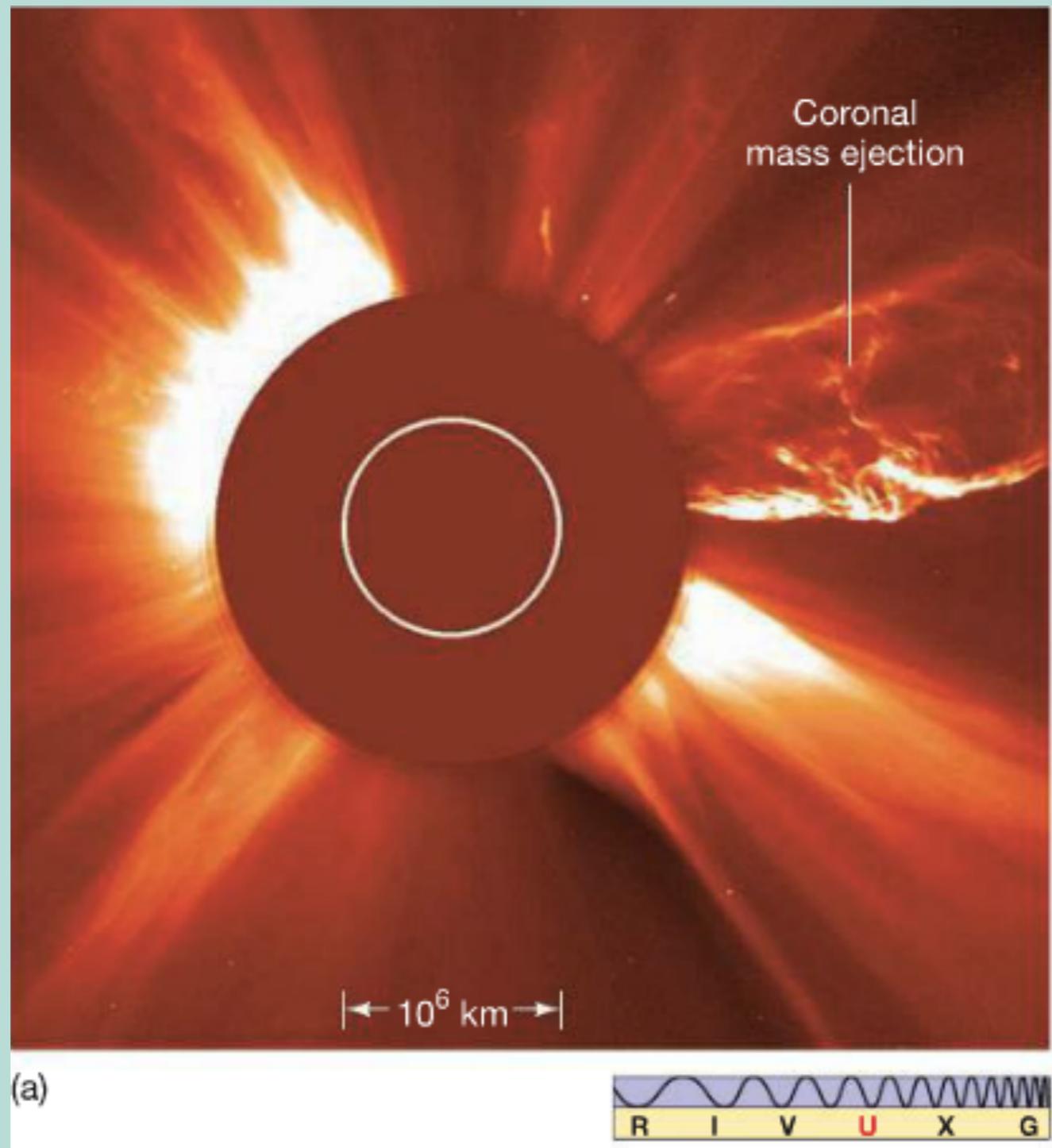
Solar Flare



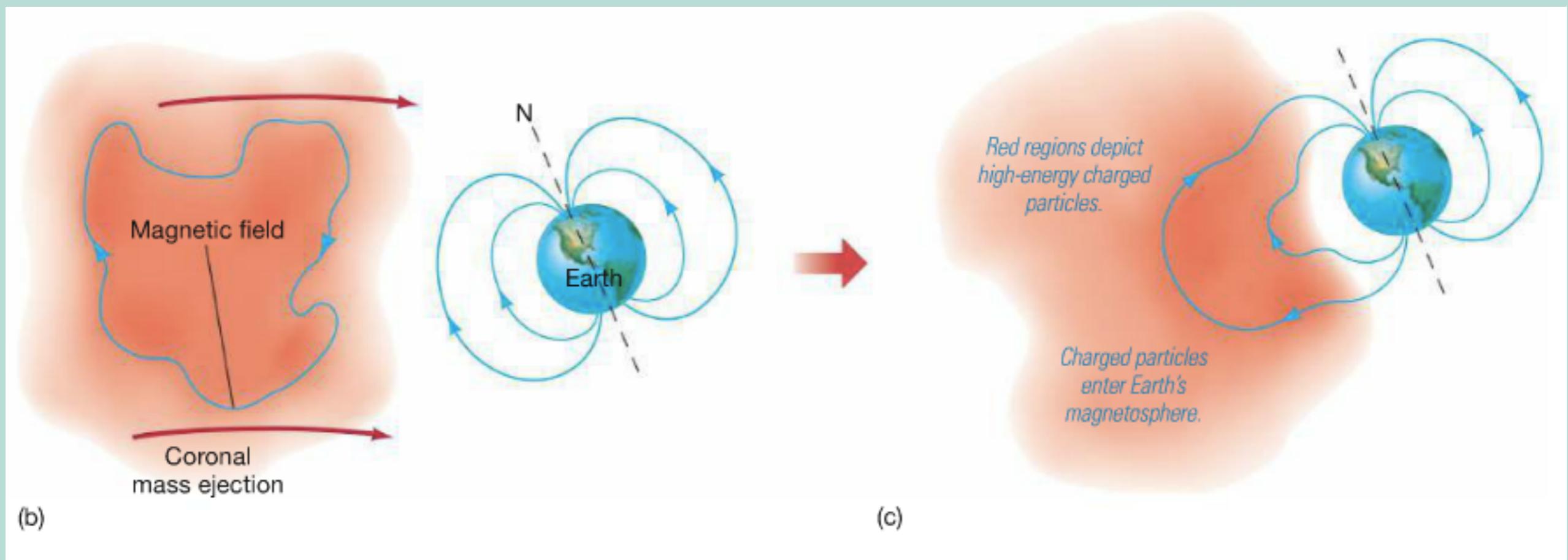
The “Carrington” Event

- September 1859
- Aurora could be seen at latitudes as far south as the Caribbean
- The Auroras in the Rockies were so bright they awoke gold miners who began preparing breakfast
- Telegraph operators could still send signals through disconnected lines!
- It has been estimated that if this happened today it would do ~ 2 trillion dollars of damage

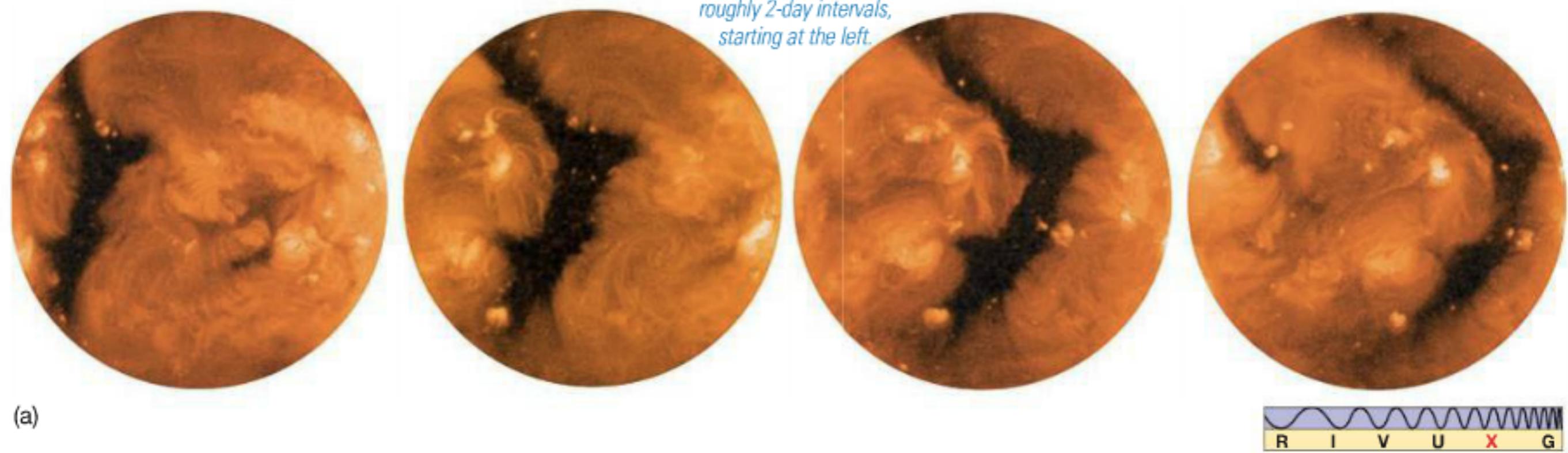
Coronal Mass Ejection (CME)



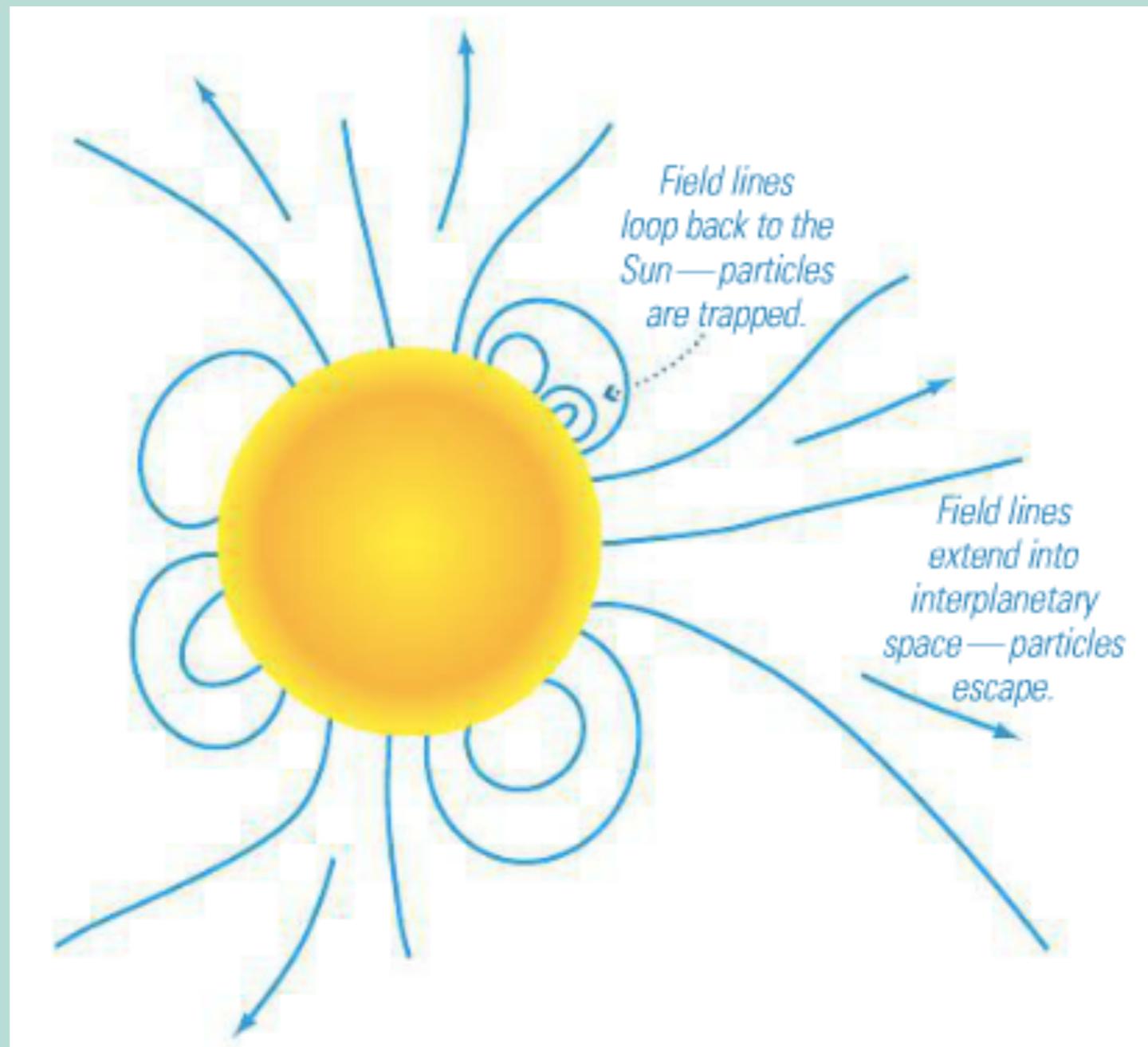
Reconnection



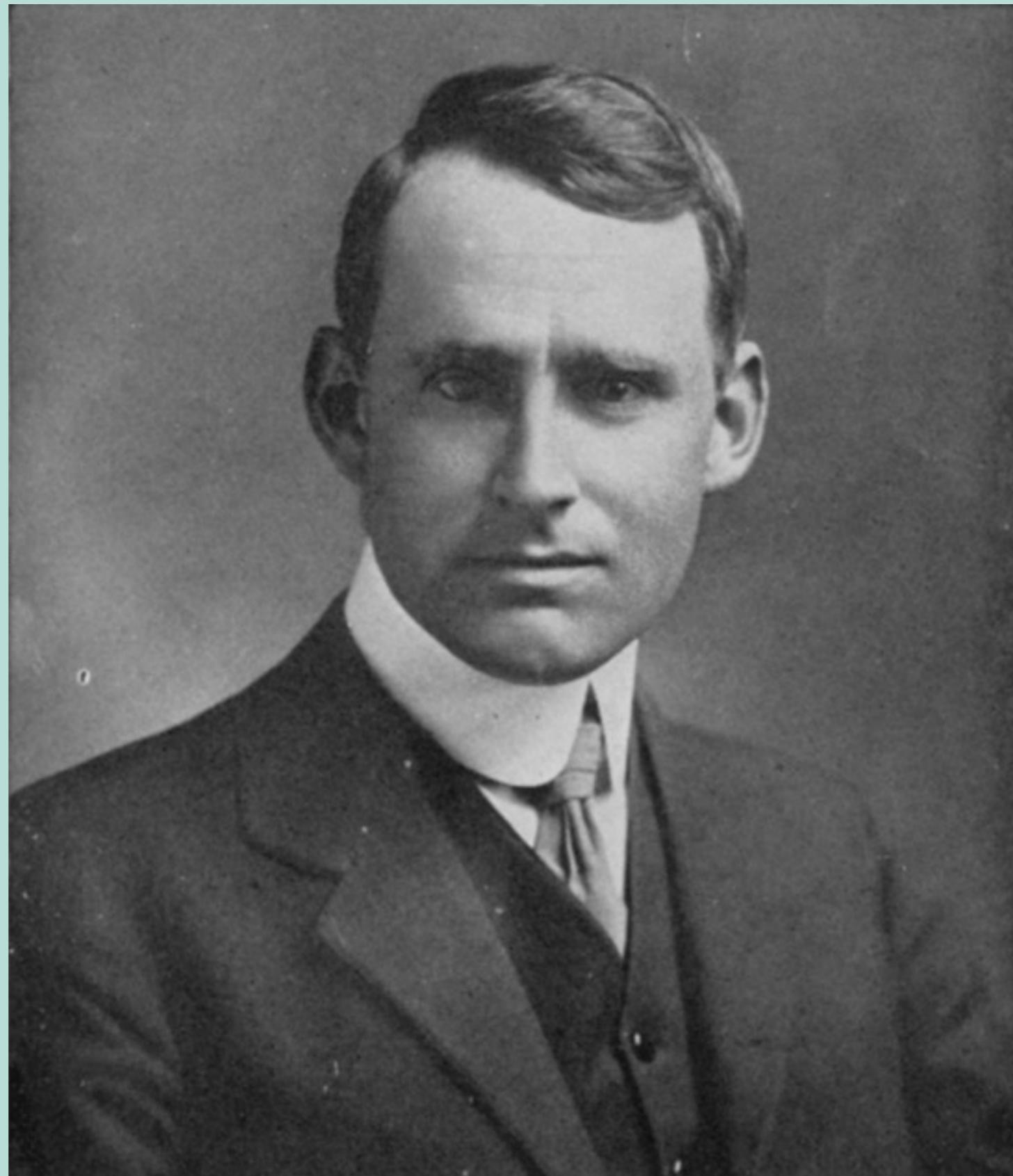
Coronal Holes



“Coronal Holes” are escaping field lines



Sir Arthur Eddington



1882 - 1944 UK

Fusion Vs Wood

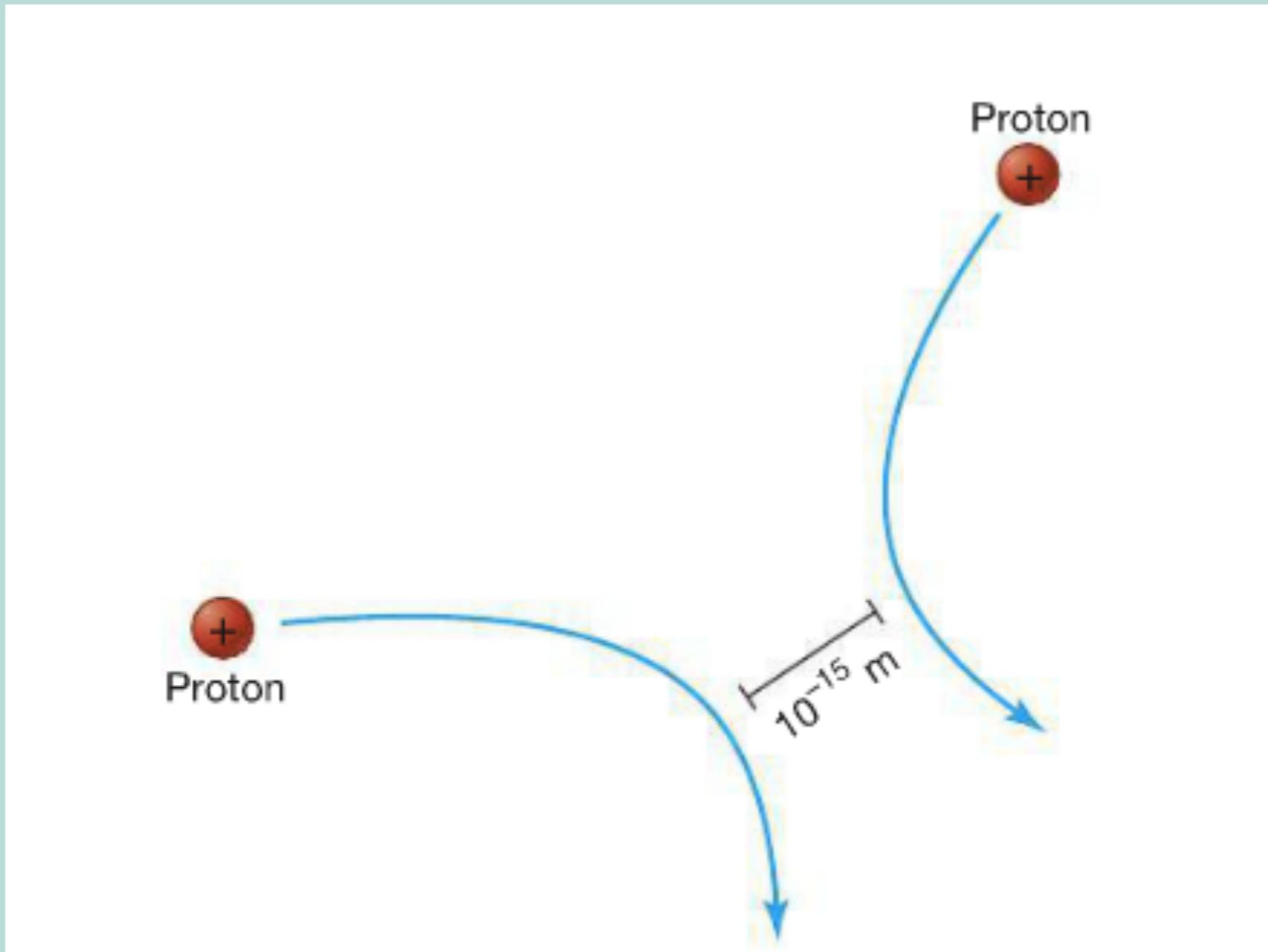
- If you divide the solar luminosity by the solar mass you get 2 mW / kilogram
- Wood generates 1,000,000 times more than that. A bon fire, merely the size of the pitiful earth would burn as bright as the sun
- How long would that fire burn though? Ans: not 5 billion years.
- Consider over 5 billion years, the amount of energy each kilogram of sun has provided so far:

10^{13} Joules!!

$$E = mc^2$$

- This means 1 kg in principle could have 10^{17} Joules packed inside!

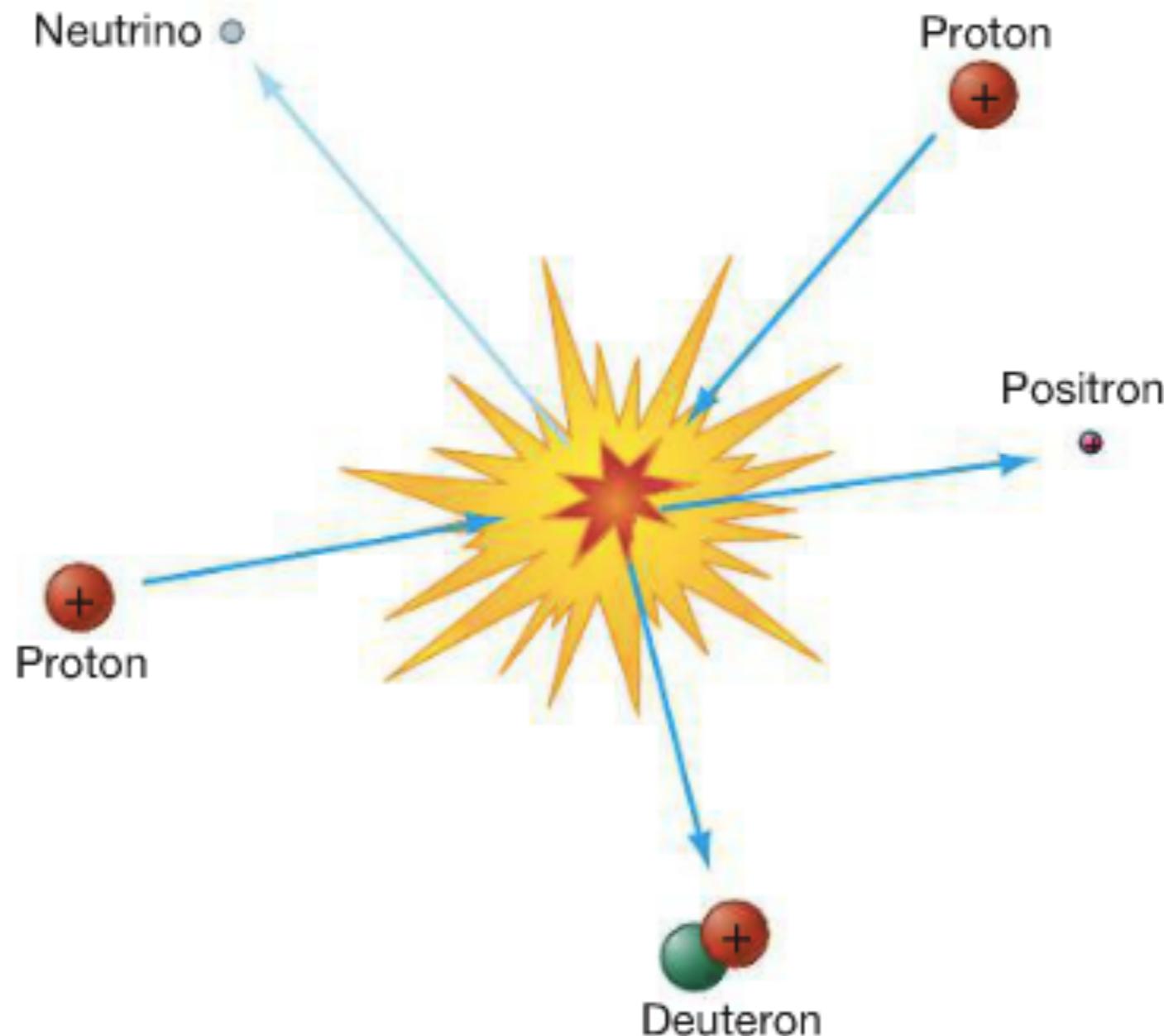
$T < 10$ million Kelvin



Electrostatic Repulsion

$T > 10$ million Kelvin

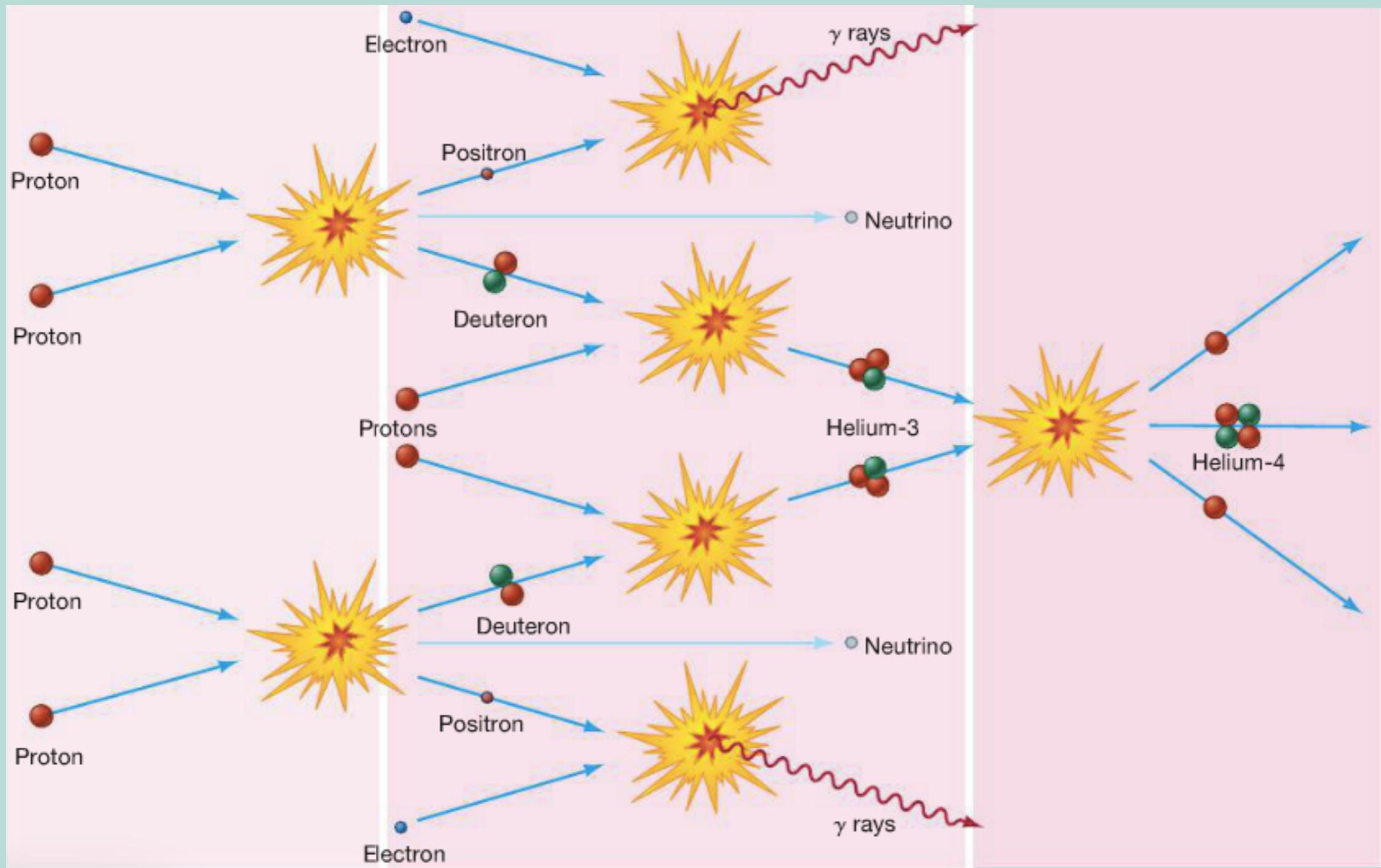
(we'll talk about these later)



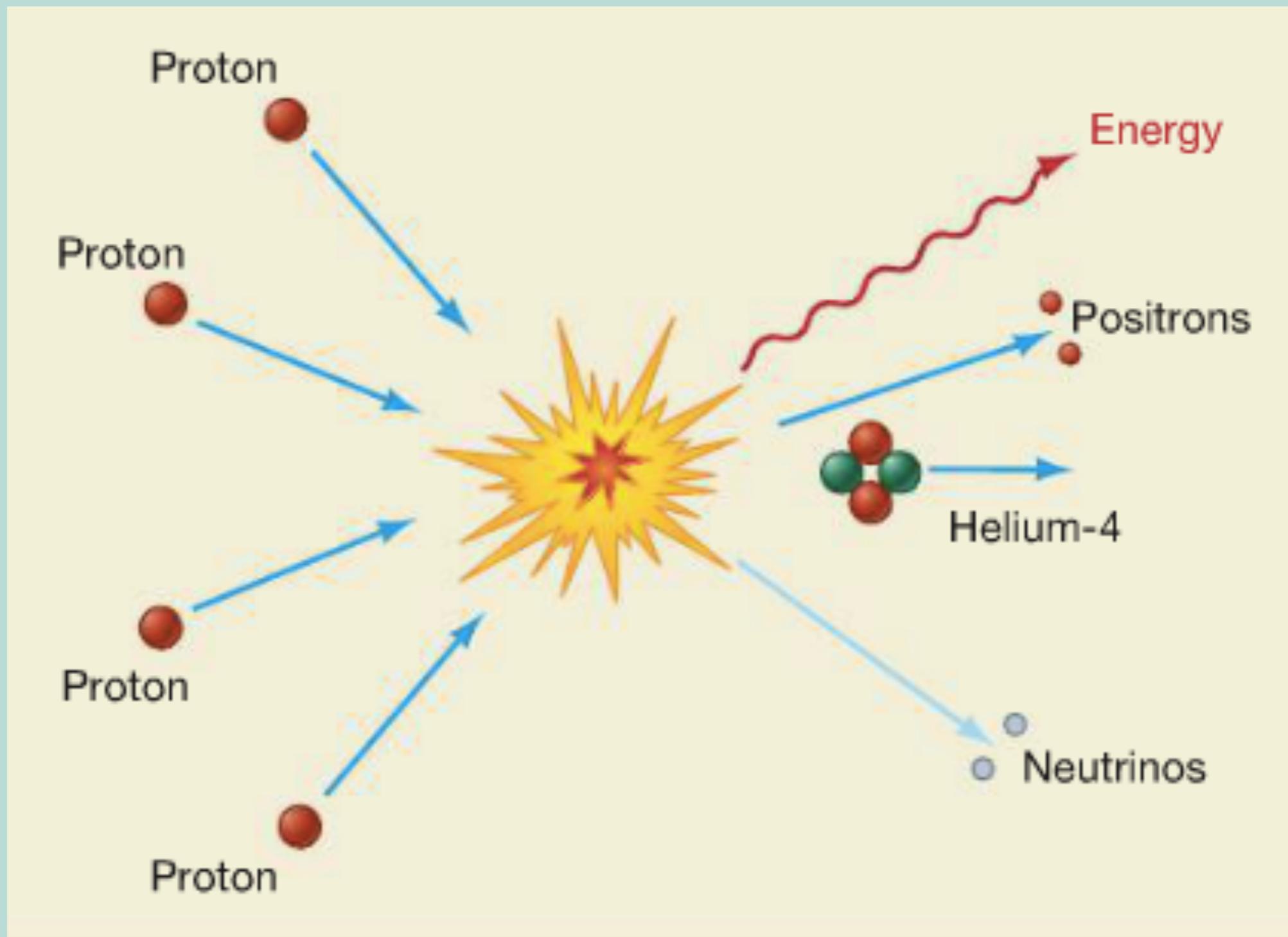
(anti-matter)

(isotopes)

How to burn 4 hydrogen nuclei

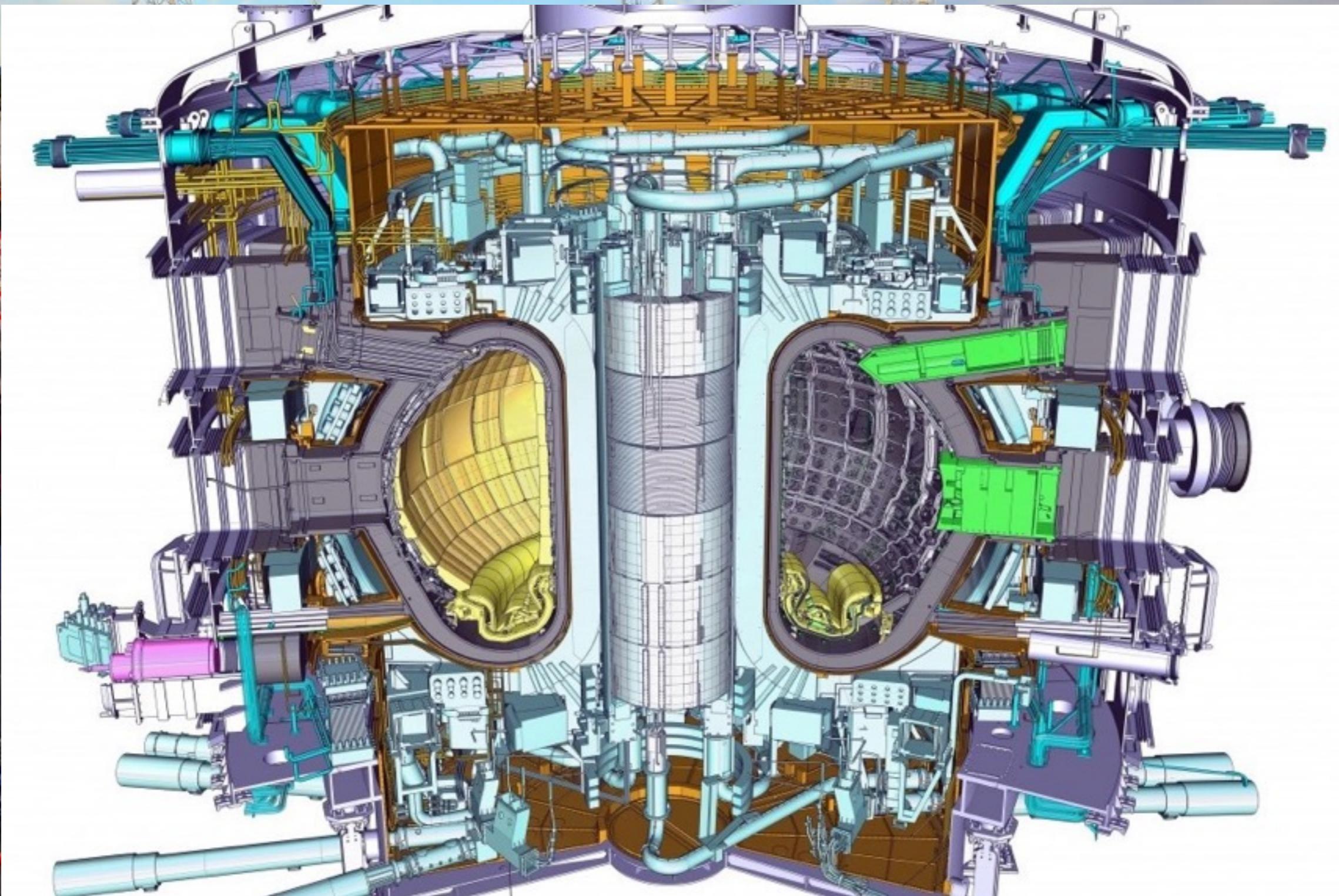


Net Fusion Reaction



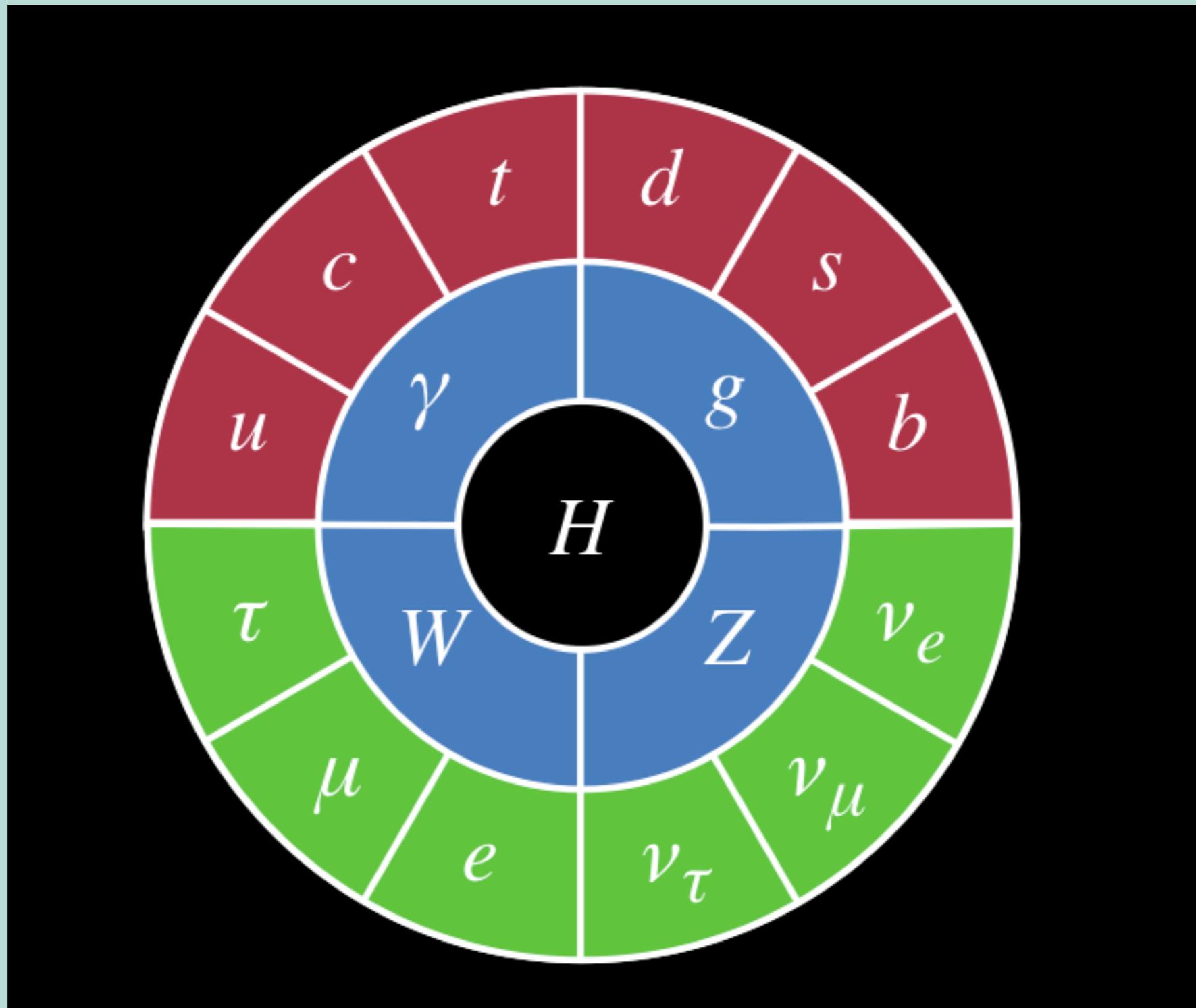
- Each fusion reaction gives you 4×10^{-12} J of energy.
- 1 kg gives you 10^{14} J of energy, so fusion (as a fraction of the total mass energy) has an efficiency of .71%
- Side note, dropping rocks into maximally rotating black holes has an efficiency of 47%
- Thus, to keep up with 4×10^{26} W of luminosity, 600 million tons of hydrogen must be fused into helium every second.
- Thus we lose 4.3 million tons of mass to pure energy per second (comparable to the solar wind).
- That might sound like a lot, but the sun is 10^{30} kg, so we can keep doing this for 5 billion years

ITER



International ThermoNuclear Reactor (2027)

Behold...

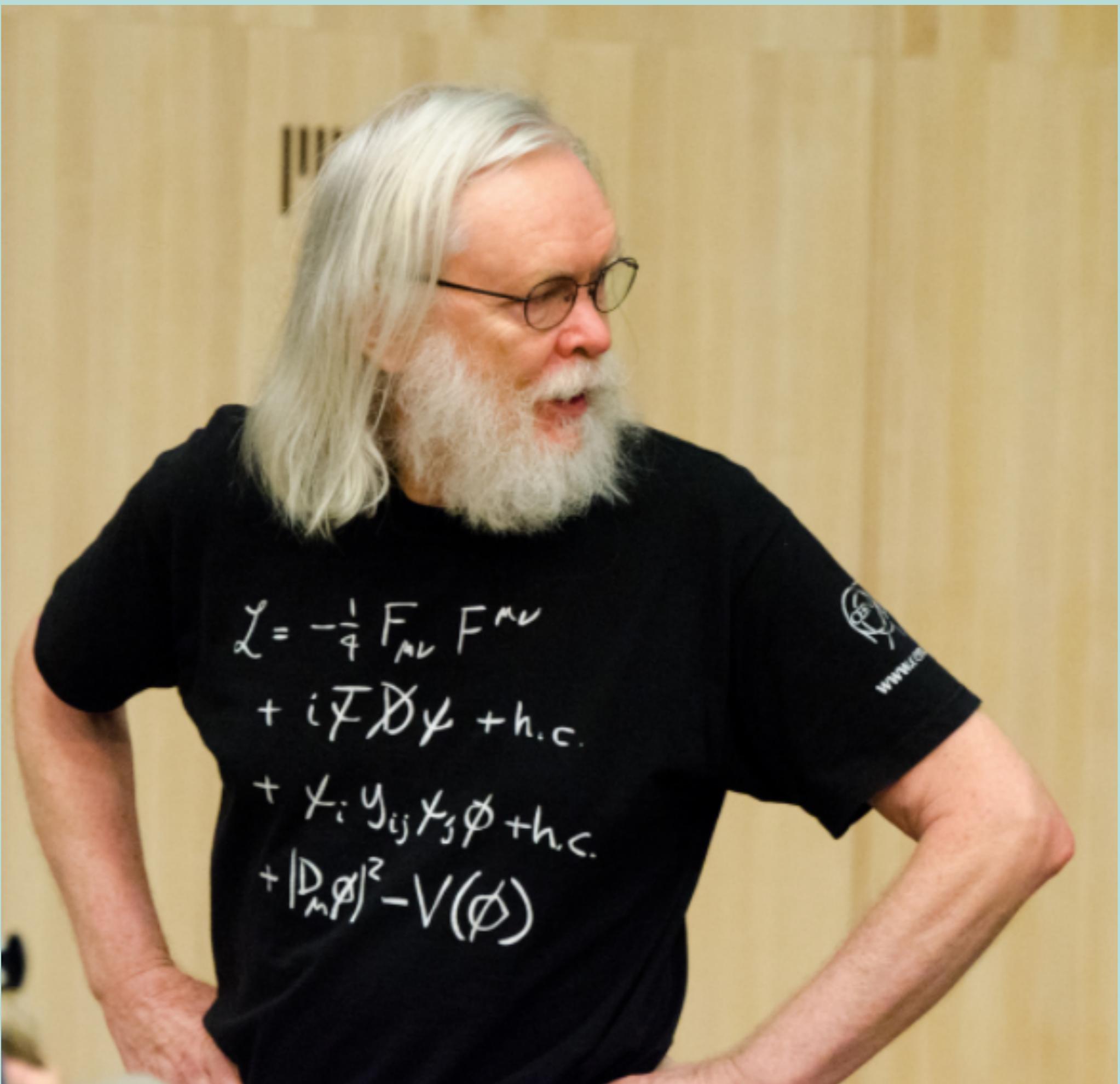


Gotta catch'em all

Quarks	Leptons		Bosons
			
Up	Down	Electron	Neutrino
			
Charm	Strange	Muon	Neutrino Muon
			
Top	Beauty	Tau	Neutrino Tau
			
Photon	Gluon	Z^0	W^-
			
		W^+	

$$\begin{aligned}
\mathcal{L}_{SM} = & -\frac{1}{2}\partial_\nu g_\mu^a \partial_\nu g_\mu^a - g_s f^{abc} \partial_\mu g_\nu^a g_\mu^b g_\nu^c - \frac{1}{4}g_e^2 f^{abc} f^{ade} g_\mu^b g_\nu^c g_\mu^d g_\nu^e - \partial_\nu W_\mu^+ \partial_\nu W_\mu^- - \\
& M^2 W_\mu^+ W_\mu^- - \frac{1}{2}\partial_\nu Z_\mu^0 \partial_\nu Z_\mu^0 - \frac{1}{2c_w^2} M^2 Z_\mu^0 Z_\mu^0 - \frac{1}{2}\partial_\mu A_\nu \partial_\mu A_\nu - ig c_w (\partial_\nu Z_\mu^0 (W_\mu^+ W_\nu^- - \\
& W_\nu^+ W_\mu^-) - Z_\nu^0 (W_\mu^+ \partial_\nu W_\mu^- - W_\mu^- \partial_\nu W_\mu^+) + Z_\mu^0 (W_\nu^+ \partial_\nu W_\mu^- - W_\nu^- \partial_\nu W_\mu^+)) - \\
& igs_w (\partial_\nu A_\mu (W_\mu^+ W_\nu^- - W_\nu^+ W_\mu^-) - A_\nu (W_\mu^+ \partial_\nu W_\mu^- - W_\mu^- \partial_\nu W_\mu^+) + A_\mu (W_\nu^+ \partial_\nu W_\mu^- - \\
& W_\nu^- \partial_\nu W_\mu^+)) - \frac{1}{2}g^2 W_\mu^+ W_\nu^- W_\nu^+ W_\mu^- + \frac{1}{2}g^2 W_\mu^+ W_\nu^- W_\mu^+ W_\nu^- + g^2 c_w^2 (Z_\mu^0 W_\mu^+ Z_\nu^0 W_\nu^- - \\
& Z_\mu^0 Z_\nu^0 W_\mu^+ W_\nu^-) + g^2 s_w^2 (A_\mu W_\mu^+ A_\nu W_\nu^- - A_\mu A_\nu W_\mu^+ W_\nu^-) + g^2 s_w c_w (A_\mu Z_\nu^0 (W_\mu^+ W_\nu^- - \\
& W_\nu^+ W_\mu^-) - 2A_\mu Z_\mu^0 W_\nu^+ W_\nu^-) - \frac{1}{2}\partial_\mu H \partial_\mu H - 2M^2 \alpha_h H^2 - \partial_\mu \phi^+ \partial_\mu \phi^- - \frac{1}{2}\partial_\mu \phi^0 \partial_\mu \phi^0 - \\
& \beta_h \left(\frac{2M^2}{g^2} + \frac{2M}{g} H + \frac{1}{2}(H^2 + \phi^0 \phi^0 + 2\phi^+ \phi^-) \right) + \frac{2M^2}{g^2} \alpha_h - \\
& g \alpha_h M (H^3 + H \phi^0 \phi^0 + 2H \phi^+ \phi^-) - \\
& \frac{1}{8}g^2 \alpha_h (H^4 + (\phi^0)^4 + 4(\phi^+ \phi^-)^2 + 4(\phi^0)^2 \phi^+ \phi^- + 4H^2 \phi^+ \phi^- + 2(\phi^0)^2 H^2) - \\
& g M W_\mu^+ W_\mu^- H - \frac{1}{2}g \frac{M}{c_w^2} Z_\mu^0 Z_\nu^0 H - \\
& \frac{1}{2}ig (W_\mu^+ (\phi^0 \partial_\mu \phi^- - \phi^- \partial_\mu \phi^0) - W_\mu^- (\phi^0 \partial_\mu \phi^+ - \phi^+ \partial_\mu \phi^0)) + \\
& \frac{1}{2}g (W_\mu^+ (H \partial_\mu \phi^- - \phi^- \partial_\mu H) + W_\mu^- (H \partial_\mu \phi^+ - \phi^+ \partial_\mu H)) + \frac{1}{2}g \frac{1}{c_w} (Z_\mu^0 (H \partial_\mu \phi^0 - \phi^0 \partial_\mu H) + \\
& M (\frac{1}{c_w} Z_\mu^0 \partial_\mu \phi^0 + W_\mu^+ \partial_\mu \phi^- + W_\mu^- \partial_\mu \phi^+) - ig \frac{s_w^2}{c_w} M Z_\mu^0 (W_\mu^+ \phi^- - W_\mu^- \phi^+) + igs_w M A_\mu (W_\mu^+ \phi^- - \\
& W_\mu^- \phi^+) - ig \frac{1-2c_w^2}{2c_w} Z_\mu^0 (\phi^+ \partial_\mu \phi^- - \phi^- \partial_\mu \phi^+) + igs_w A_\mu (\phi^+ \partial_\mu \phi^- - \phi^- \partial_\mu \phi^+) - \\
& \frac{1}{4}g^2 W_\mu^+ W_\mu^- (H^2 + (\phi^0)^2 + 2\phi^+ \phi^-) - \frac{1}{8}g^2 \frac{1}{c_w^2} Z_\mu^0 Z_\nu^0 (H^2 + (\phi^0)^2 + 2(2s_w^2 - 1)^2 \phi^+ \phi^-) - \\
& \frac{1}{2}g^2 \frac{s_w^2}{c_w} Z_\mu^0 \phi^0 (W_\mu^+ \phi^- + W_\mu^- \phi^+) - \frac{1}{2}ig^2 \frac{s_w^2}{c_w} Z_\mu^0 H (W_\mu^+ \phi^- - W_\mu^- \phi^+) + \frac{1}{2}g^2 s_w A_\mu \phi^0 (W_\mu^+ \phi^- + \\
& W_\mu^- \phi^+) + \frac{1}{2}ig^2 s_w A_\mu H (W_\mu^+ \phi^- - W_\mu^- \phi^+) - g^2 \frac{s_w}{c_w} (2c_w^2 - 1) Z_\mu^0 A_\mu \phi^+ \phi^- - \\
& g^2 s_w^2 A_\mu A_\mu \phi^+ \phi^- + \frac{1}{2}ig_s \lambda_{ij}^\alpha (q_i^\sigma \gamma^\mu q_j^\sigma) g_\mu^\alpha - \bar{e}^\lambda (\gamma \partial + m_e^\lambda) e^\lambda - \bar{\nu}^\lambda (\gamma \partial + m_\nu^\lambda) \nu^\lambda - \bar{u}_j^\lambda (\gamma \partial + \\
& m_u^\lambda) u_j^\lambda - \bar{d}_j^\lambda (\gamma \partial + m_d^\lambda) d_j^\lambda + igs_w A_\mu \left(-(\bar{e}^\lambda \gamma^\mu e^\lambda) + \frac{2}{3} (\bar{u}_j^\lambda \gamma^\mu u_j^\lambda) - \frac{1}{3} (\bar{d}_j^\lambda \gamma^\mu d_j^\lambda) \right) + \\
& \frac{ig}{4c_w} Z_\mu^0 \left((\bar{\nu}^\lambda \gamma^\mu (1 + \gamma^5) \nu^\lambda) + (\bar{e}^\lambda \gamma^\mu (4s_w^2 - 1 - \gamma^5) e^\lambda) + (\bar{d}_j^\lambda \gamma^\mu (\frac{4}{3}s_w^2 - 1 - \gamma^5) d_j^\lambda) + \right. \\
& \left. (\bar{u}_j^\lambda \gamma^\mu (1 - \frac{8}{3}s_w^2 + \gamma^5) u_j^\lambda) \right) + \frac{ig}{2\sqrt{2}} W_\mu^+ \left((\bar{\nu}^\lambda \gamma^\mu (1 + \gamma^5) U^{lep}{}^\lambda{}_\kappa e^\kappa) + (\bar{u}_j^\lambda \gamma^\mu (1 + \gamma^5) C_{\lambda\kappa} d_j^\kappa) \right) + \\
& \frac{ig}{2\sqrt{2}} W_\mu^- \left((\bar{e}^\kappa U^{lep\dagger}{}_{\kappa\lambda} \gamma^\mu (1 + \gamma^5) \nu^\lambda) + (\bar{d}_j^\kappa C_{\kappa\lambda}^\dagger \gamma^\mu (1 + \gamma^5) u_j^\lambda) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^+ \left(-m_e^\kappa (\bar{\nu}^\lambda U^{lep}{}_{\lambda\kappa} (1 - \gamma^5) e^\kappa) + m_\nu^\lambda (\bar{\nu}^\lambda U^{lep}{}_{\lambda\kappa} (1 + \gamma^5) e^\kappa) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^- \left(m_e^\lambda (\bar{e}^\lambda U^{lep\dagger}{}_{\lambda\kappa} (1 + \gamma^5) \nu^\kappa) - m_\nu^\kappa (\bar{e}^\lambda U^{lep\dagger}{}_{\lambda\kappa} (1 - \gamma^5) \nu^\kappa) \right) - \frac{g m_e^\lambda}{2M} H (\bar{\nu}^\lambda \nu^\lambda) - \\
& \frac{g m_e^\lambda}{2M} H (\bar{e}^\lambda e^\lambda) + \frac{ig}{2} \frac{m_e^\lambda}{M} \phi^0 (\bar{\nu}^\lambda \gamma^5 \nu^\lambda) - \frac{ig}{2} \frac{m_e^\lambda}{M} \phi^0 (\bar{e}^\lambda \gamma^5 e^\lambda) - \frac{1}{4} \bar{\nu}_\lambda M_{\lambda\kappa}^R (1 - \gamma_5) \bar{\nu}_\kappa - \\
& \frac{1}{4} \bar{\nu}_\lambda M_{\lambda\kappa}^R (1 - \gamma_5) \bar{\nu}_\kappa + \frac{ig}{2M\sqrt{2}} \phi^+ \left(-m_d^\kappa (\bar{u}_j^\lambda C_{\lambda\kappa} (1 - \gamma^5) d_j^\kappa) + m_u^\lambda (\bar{u}_j^\lambda C_{\lambda\kappa} (1 + \gamma^5) d_j^\kappa) \right) + \\
& \frac{ig}{2M\sqrt{2}} \phi^- \left(m_d^\lambda (\bar{d}_j^\lambda C_{\lambda\kappa}^\dagger (1 + \gamma^5) u_j^\kappa) - m_u^\kappa (\bar{d}_j^\lambda C_{\lambda\kappa}^\dagger (1 - \gamma^5) u_j^\kappa) \right) - \frac{g m_e^\lambda}{2M} H (\bar{u}_j^\lambda u_j^\lambda) - \\
& \frac{g m_d^\lambda}{2M} H (\bar{d}_j^\lambda d_j^\lambda) + \frac{ig}{2} \frac{m_d^\lambda}{M} \phi^0 (\bar{u}_j^\lambda \gamma^5 u_j^\lambda) - \frac{ig}{2} \frac{m_d^\lambda}{M} \phi^0 (\bar{d}_j^\lambda \gamma^5 d_j^\lambda) + \bar{G}^a \partial^2 G^a + g_s f^{abc} \partial_\mu \bar{G}^a G^b g_\mu^c + \\
& \bar{X}^+ (\partial^2 - M^2) X^+ + \bar{X}^- (\partial^2 - M^2) X^- + \bar{X}^0 (\partial^2 - \frac{M^2}{c_w^2}) X^0 + \bar{Y} \partial^2 Y + ig c_w W_\mu^+ (\partial_\mu \bar{X}^0 X^- - \\
& \partial_\mu \bar{X}^+ X^0) + igs_w W_\mu^+ (\partial_\mu \bar{Y} X^- - \partial_\mu \bar{X}^+ Y) + ig c_w W_\mu^- (\partial_\mu \bar{X}^- X^0 - \\
& \partial_\mu \bar{X}^0 X^+) + igs_w W_\mu^- (\partial_\mu \bar{X}^- Y - \partial_\mu \bar{Y} X^+) + ig c_w Z_\mu^0 (\partial_\mu \bar{X}^+ X^+ - \\
& \partial_\mu \bar{X}^- X^-) + igs_w A_\mu (\partial_\mu \bar{X}^+ X^+ - \\
& \partial_\mu \bar{X}^- X^-) - \frac{1}{2}g M \left(\bar{X}^+ X^+ H + \bar{X}^- X^- H + \frac{1}{c_w^2} \bar{X}^0 X^0 H \right) + \frac{1-2c_w^2}{2c_w} ig M (\bar{X}^+ X^0 \phi^+ - \bar{X}^- X^0 \phi^-) + \\
& \frac{1}{2c_w} ig M (\bar{X}^0 X^- \phi^+ - \bar{X}^0 X^+ \phi^-) + ig M s_w (\bar{X}^0 X^- \phi^+ - \bar{X}^0 X^+ \phi^-) + \\
& \frac{1}{2}ig M (\bar{X}^+ X^+ \phi^0 - \bar{X}^- X^- \phi^0) .
\end{aligned}$$

75% of Physics.



Solar Neutrino Problem

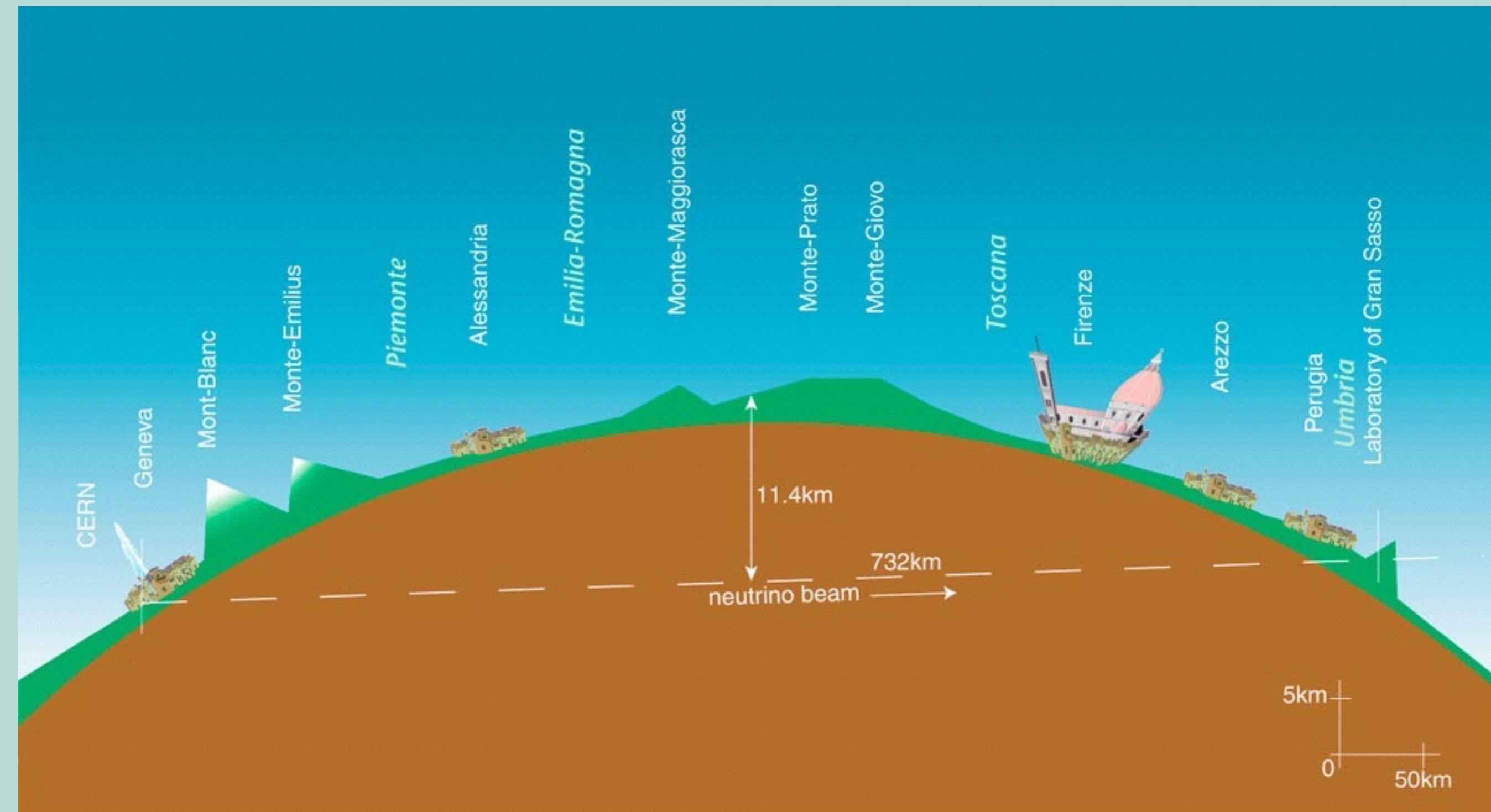
- With the amount of data that we have on the sun, we know exactly how many neutrinos are created in the core
- We also know how likely those neutrinos are to interact with our detectors
- We see a mismatch. This was explained by “neutrino oscillations”, experimentally verified in 2002



FACT: about 65 million neutrinos pass through your thumbnail every second.

Learn Something
New Every Day
LSNED.com

The Earth is transparent to neutrinos.





IceCube Laboratory

Data is collected here and sent by satellite to the data warehouse at UW–Madison



Digital Optical Module (DOM)

5,160 DOMs deployed in the ice

50 m

1450 m

2450 m

IceTop

86 strings of DOMs,
set 125 meters apart

IceCube
detector

DeepCore

Antarctic bedrock

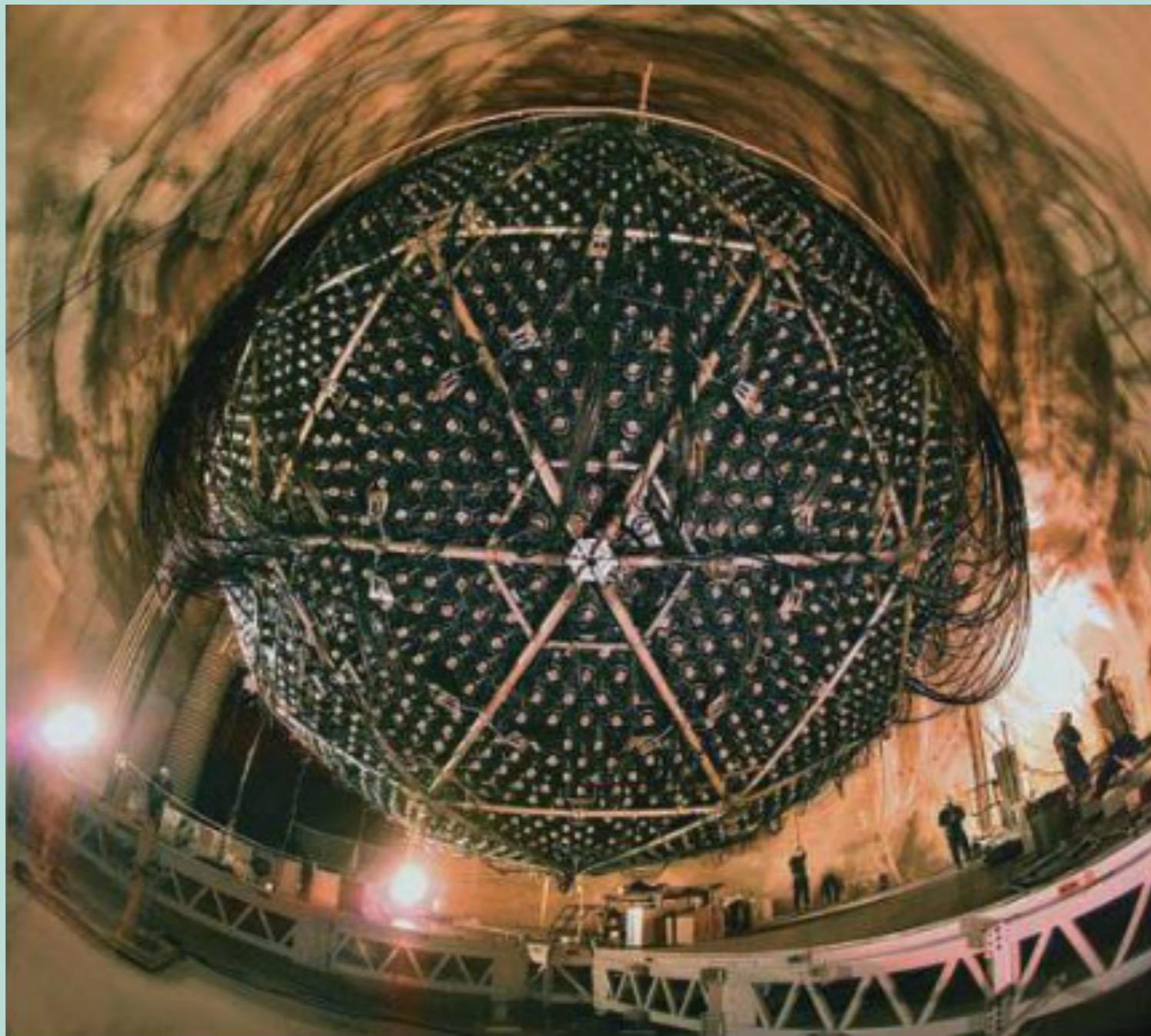
Amundsen–Scott South Pole Station, Antarctica
A National Science Foundation-managed research facility

60 DOMs
on each string

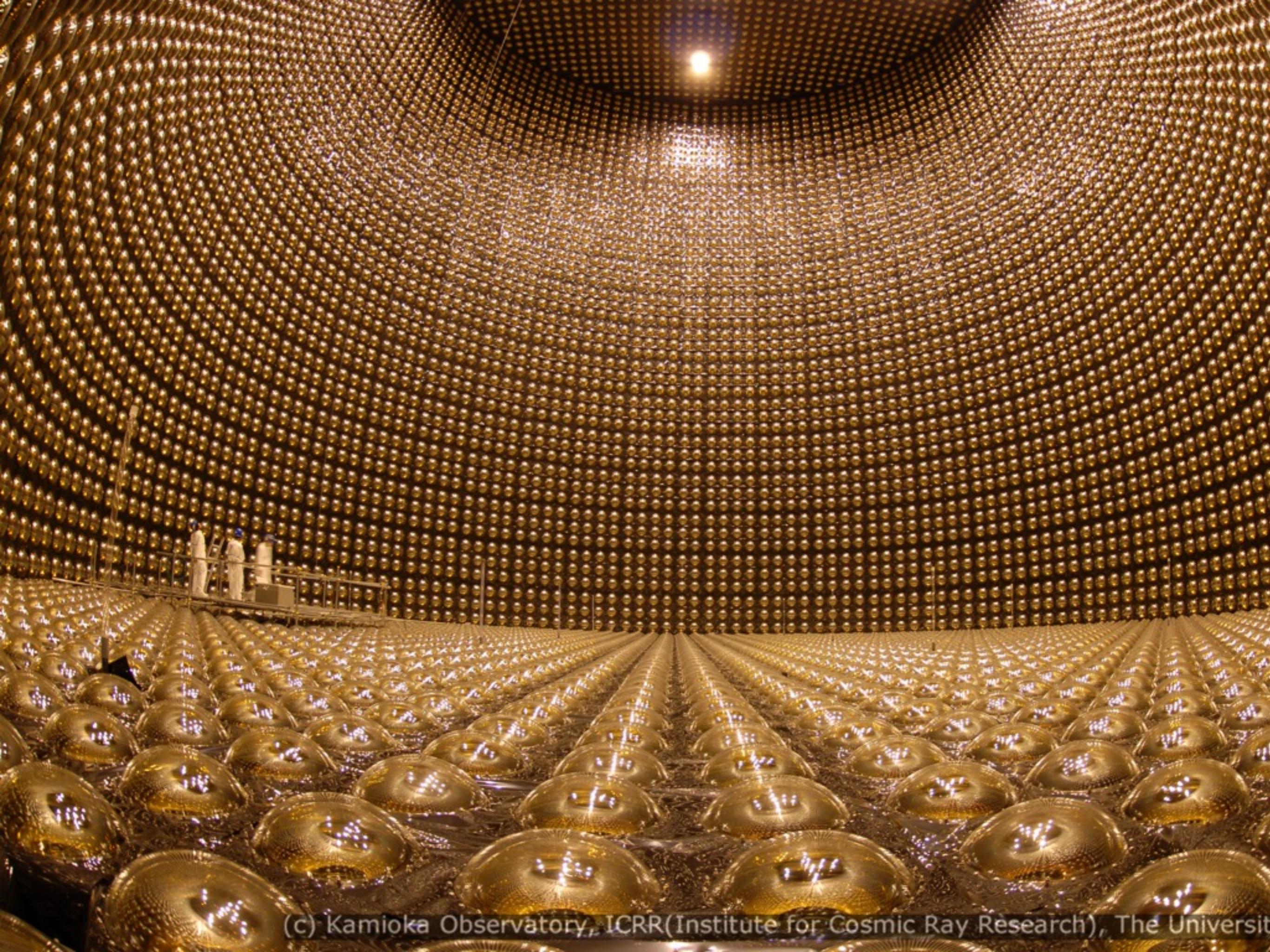
DOMs
are 17
meters
apart



Ontario, CA



Sudbury Neutrino Solar Laboratory



(c) Kamioka Observatory, ICRR (Institute for Cosmic Ray Research), The University