**Original text:** Raiter, Schaerer, and Fosbury (2013) discuss the equivalent widths of recombination lines Lyα and He II λ1640 in a larger critique of the assumptions made with Case B approximations. They discuss the dependence of Lyα on temperature. Since the Lyα luminosity increases more rapidly than the continuum flux near Lyα, the equivalent width is sensitive to temperature. Alternatively, the He II λ1640 equivalent width depends more strongly on ionization parameter since the line luminosity changes due to the “Stasinska-Tylenda effect.” With an ionization energy 4 times larger than Lyα, He II λ1640 emission is much less abundant on our grids. Lyα emits strongly on our plane with typical emission around log(*W*λ) = 3.5 (Figure 4a, row b), while He II λ1640 emits faintly across our grid with typical emission around log(*W*λ) = 1.0 (Figure 4a, row e).

**Replacement text:** Raiter, Schaerer, and Fosbury (2013) [I think this should be the same authors but in 2010, but the references section gives a different paper that I can’t find] show that at very low metallicities (< 0.01 Z\_solar) emission line predictions for Lyα and He II λ1640 begin to deviate from their case B values due to enhanced collisional effects from a lack of coolants. In hydrogen, this leads to increase in *n* = 2 populations, from which additional ionization can occur, thus increasing the overall ionization rate. In most physical situations, including the ones presented in this paper, the case B approximation is appropriate given that our low metallicity simulations still contain a significant number of coolants, so collisional effects are less important than one would expect for a primordial composition nebula. However, we note very low metallicity environments can occur in high-*z* galaxies, so observers should remain cautious about using common approximations.

[I THINK THIS PARAGRAPH SHOULD BE MOVED TO THE METALLICITY SECTION]