

# CMB polarization with the Bicep and Keck experiments

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## 1 Introduction

### 1.1 A Brief History of Inflation

The modern study of cosmology began in 1915 with Einstein's development of the general theory of Relativity [6]. By combining the Einstein Field Equations with the assumptions of a homogeneous and isotropic universe, Friedmann developed a cosmological model which forms the core of the standard model of cosmology[7]. Hubble's observations in the 1930s showed an expanding universe[9] providing the first evidence for a Big Bang origin. This was further supported by the detection of the Cosmic Microwave Background by Penzias and Wilson in 1965[12].

Increasing experimental precision revealed a uniform CMB and isotropic universe on scales exceeding the size of the horizon predicted by Friedmann cosmological models. In 1981 Guth suggested a small period of exponential expansion to solve these two problems. This theory of inflation has since seen further success in explaining the growth of cosmic structures and the lack of observed magnetic monopoles but has been criticized due to its need for fine-tuned initial conditions.

## 1.2 Origins of the CMB

One of the most successful predictions of inflation is the presence of a near uniform isotropic photon source known as the Cosmic Microwave Background (CMB). The CMB has the near uniform temperature spectrum of a 2.73K blackbody and is the source of the oldest photons in the universe. In the hot big bang, the early universe was an energetic plasma in which photons scattered continuously from free protons and electrons which prevented propagation over large distances. As the universe expanded, the temperature of this plasma decreased adiabatically until the free electrons and protons could combine to form neutral Hydrogen. After this era of recombination about 380,000 years after the big bang, the photons decoupled from the plasma and began free streaming through the universe, redshifting due to continued expansion. We observe these photons as a 2D "surface of last scattering".

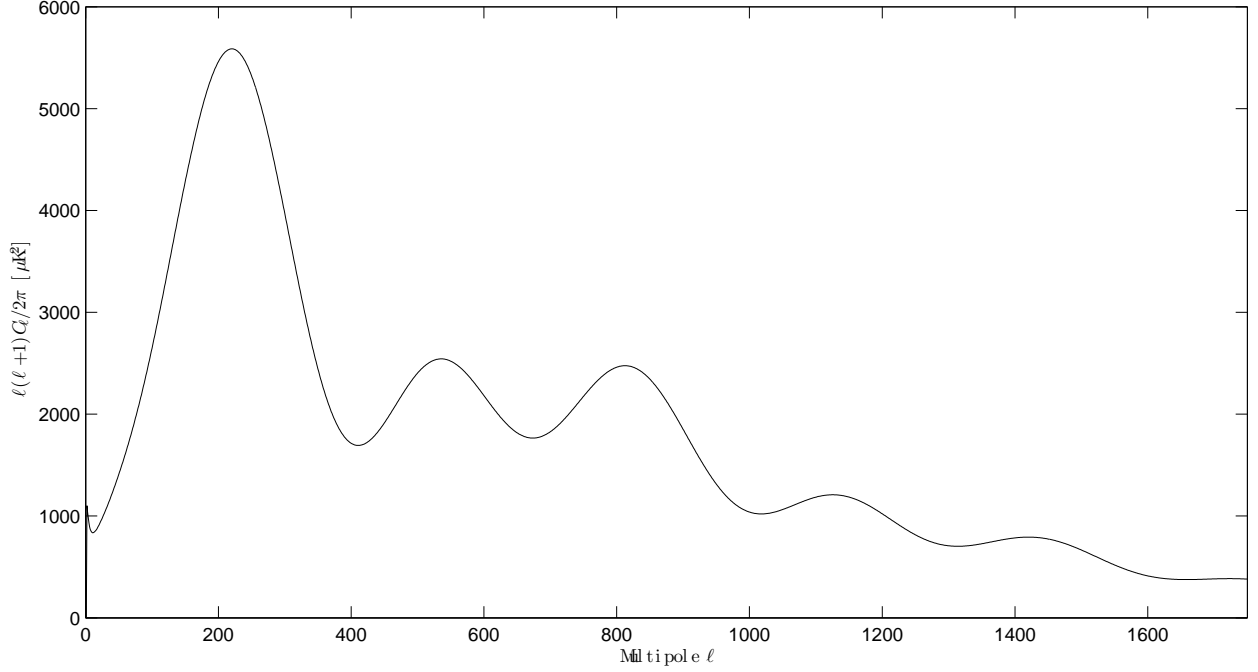
## 2 Science with the CMB

### 2.1 Mapping the CMB

Precision cosmology experiments have shown that the CMB is not truly uniform. Results from COBE in 1992[4] showed the CMB to have a nearly isotropic temperature spectrum corresponding to a 2.73K blackbody but with fluctuations on the order of  $10 \mu\text{K}$ . We examine these temperature perturbations with the spherical harmonics:

$$\delta T(\theta, \phi) = \sum_{m=-\ell}^{\ell} \sum_{\ell=0}^{\infty} a_{\ell m} Y_{\ell m}(\theta, \phi) \quad (1)$$

Although this decomposition is a powerful tool, the universe is isotropic only in a statistical sense and we therefore cannot predict any individual  $a_{\ell m}$ . We instead take an average over  $m$  for a given  $\ell$  to form the angular power spectrum:



**Figure 1:** The temperature power spectrum of the CMB. The primary and acoustic peaks can be used to constrain cosmological parameters that describe the geometry, evolution and composition of our universe.

$$C_\ell = \frac{1}{2\ell + 1} \sum_{m=-\ell}^{\ell} |a_{\ell m}|^2 \quad (2)$$

The temperature power spectrum shown in Figure 1 contains much of the information concerning the geometry, evolution, and composition of our universe. The simple  $\Lambda$ CDM model is based on the six parameters shown in Table 1.

## 2.2 The 6 Parameter Model

The first three parameters in Table 1 -  $h^2\Omega_b$ ,  $h^2\Omega_c$ , and  $\Omega_\Lambda$  - describe relative contribution to the composition of our universe from baryons, cold dark matter, and dark energy respectively. The matter to baryon ratio is related to the size of the first three peaks of the Temperature power spectrum in Figure 1. The scalar spectral index  $n_s$  measures the deviation of the power spectrum from the scale-invariant case. The curvature fluctuation amplitude  $\Delta_{\mathcal{R}}^2$  is a

| Parameter                                     | Symbol                   | Value                 |
|---|--------------------------|-----------------------|
| Baryon Density                                | $h^2\Omega_b$            | 0.02240               |
| Cold Dark Matter Density                      | $h^2\Omega_m$            | 0.1146                |
| Dark Energy Density ( $w = -1$ )              | $\Omega_\Lambda$         | 0.7181                |
| Scalar Spectral Index                         | $n_s$                    | 0.9646                |
| Curvature Fluctuation Amplitude $k_0 = 0.002$ | $\Delta_{\mathcal{R}}^2$ | $2.43 \times 10^{-9}$ |
| Reionization Optical Depth                    | $\tau$                   | 0.0800                |

**Table 1:** The six parameters in the simple  $\Lambda$ CDM cosmological model as shown in the WMAP 9 year results[8]

measure of curvature fluctuation at different angular scales. Finally, the reionization optical depth can be used to estimate the damping of CMB fluctuations due to the reionization of cosmic baryons which turns the universe more opaque for CMB photons.

These 6 parameters are largely degenerate with each other in CMB studies alone. To break the degeneracies, CMB data is often combined with data from other cosmological studies which can set constraints independent of the CMB. Studies of Type Ia supernovae use well characterized spectral emissions to create a cosmic distance ladder which can constrain the Hubble constant and place independent constraints on the composition of the universe. An independent measure of these parameters can be made by studying Baryon Acoustic Oscillations. The attraction of matter by a primordial overdensity leads to increased heat and a large outward pressure. These pressure waves of baryons were left behind at fixed radii after photon decoupling leaving a large overdensity at the sound horizon. By comparing the overdensities in the CMB with the distribution of modern galaxies, BAO places constraints on cosmic expansion independent from those set by CMB or supernovae studies.

## 2.3 Polarization Anisotropies

In addition to the temperature anisotropies CMB photons are partially polarized due to Thompson scattering in a non-uniform temperature field. This CMB polarization was predicted at the  $\approx 10\%$  level by Bond et al [5] and first detected by DASI in 2002 [11]. Polnarev

[15] realized that an inflationary period would imprint an additional polarization signal on the CMB due to the production of gravitational waves.

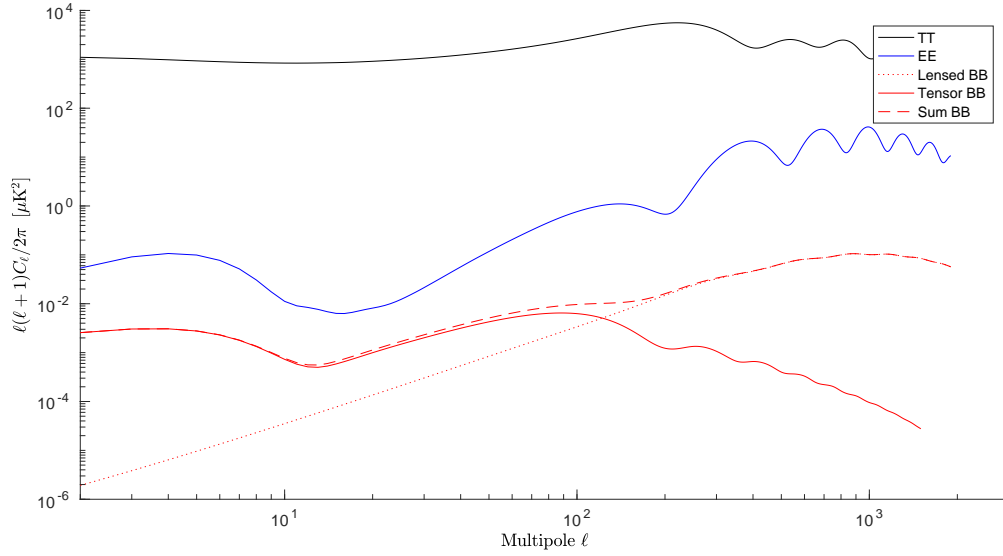
This inflationary gravitational wave (IGW) signal is expected to be orders of magnitude fainter than the signal due to Thompson scattering and difficult to separate in the Stokes parameter space conventionally used when studying linearly polarized light. It was shown by Kamionkowski et al [10] that the polarization could instead be parameterized by a gradient (E-mode) and curl (B-mode) component. The signal produced by Thompson scattering occurs in the presence of a temperature quadrupole and follows the temperature gradients of the hot plasma at recombination and can therefore produce only E-modes. The tensor perturbations due to passing gravitational waves are not restricted in this way and can thus generate both E and B modes. The signal produced by these two methods is characterised by the tensor to scalar ratio  $r$ .

In addition to the temperature power spectrum, we can then examine the E, B power spectrum as well as cross correlations between the temperature and polarization power spectra. Figure 2 shows the theoretical power spectra of CMB perturbations in the  $\Lambda$ CDM standard model of Cosmology with the addition of an  $r = 0.1$  IGW B-mode signal.

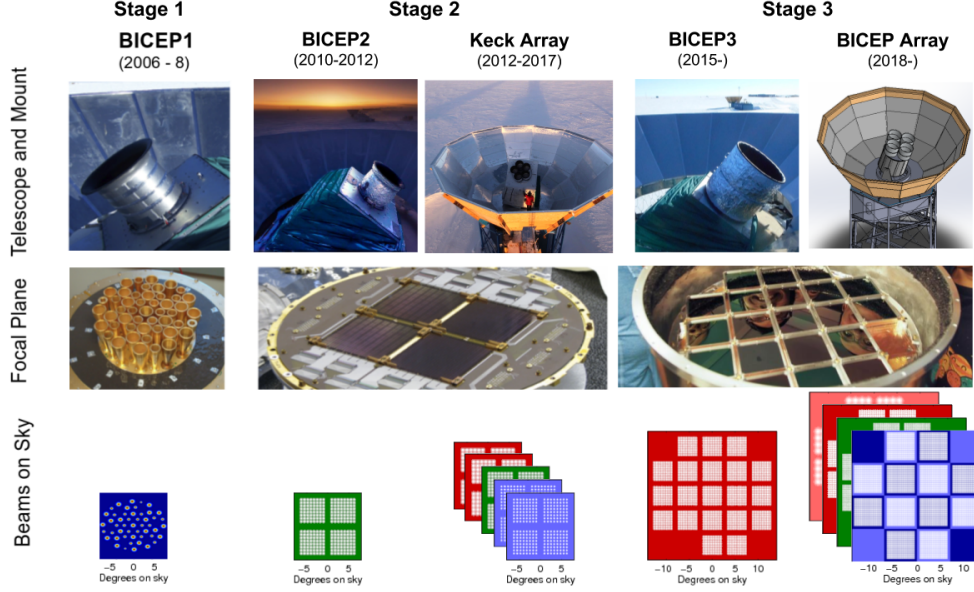
### 3 The Bicep and Keck Program

The Bicep/Keck experiments are a staged series of small aperture ground based telescopes which aim to produce extremely deep degree-scale polarization maps of the CMB. Each generation of receiver builds on the experience gained from the previous generation while pushing deeper in sensitivity. This progression is shown in Figure 3.

Bicep1 was deployed to the south pole in 2006 and used 98 feedhorn coupled bolometers observing at 100GHz and 150GHz. Over three observing seasons the strategies developed for observation, calibration and systematics control proved the efficacy of small aperture



**Figure 2:** The CMB power spectrum of the  $\Lambda$ CDM standard model of cosmology as calculated by CAMB using parameters from Planck 2013. The temperature anisotropies (black) contain orders of magnitude more power than the E-mode (blue) and B-mode (red) polarization anisotropies. The BB spectrum is split into a lensing component (dotted) and a tensor component (solid) plotted at the  $r = 0.1$  level.



**Figure 3:** The progression of the Bicep/Keck program from Bicep1 to Bicep Array. The bottom row shows the beam patterns of the focal planes on the sky. With the exception of Bicep1 the focal plane colors correspond to band centers of: 35GHz - pink, 95GHz - red, 150GHz - green, 220 GHz - light blue, 270GHz - dark blue.

refractors for CMB polarization studies and established the leading upper bounds on inflation at  $r < 0.70$  [1].

Building on the techniques developed with its predecessor, Bicep2 replaced Bicep1 in 2009. It exchanged feedhorn coupled bolometers for antenna-coupled transition edge sensor (TES) bolometer arrays developed at JPL which have been used in every subsequent telescope in the series. Concentrating its observing power at 150GHz, in March 2014 Bicep2 announced a detection of excess signal in its observing band consistent with an IGW signal of  $r = 0.2$ [2]. However, the interpretation of this excess as IGW signal relied heavily on models of polarized dust emission which had not been highly constrained at the time. Later that year new high frequency maps from the Planck experiment indicated that the dust models had underestimated polarized emission in the faintest sky regions [13]. A joint analysis with Planck and cross correlation between the Planck 353GHz and Bicep2 150GHz maps showed that a substantial part of the observed excess in Bicep2 was due to polarized dust emissions [3]. This joint analysis established a new upper limit of  $r < 0.12$

The Keck Array deployed to the south pole in 2012 with five 150GHz receivers similar to Bicep2. These additional receivers confirmed the excess signal found with Bicep2 and contributed to the March 2014 results. In addition to extending the Bicep2 survey depth at 150GHz the Keck array has extended observations into three other observing bands at 95GHz, 220GHz and 270GHz. The extension into other frequencies harnesses the Bicep/Keck program's proven capability to make deep maps to further constrain galactic foregrounds and refine the models of polarized dust emission. The Keck array is in its final observing season with four receivers in the 220GHz band and one at 270GHz.

In the fall of 2014 Bicep3 was deployed at the south pole to run concurrently with the Keck array. Bicep3 vastly expands the design of the Bicep2 instrument with a focal plane containing 2500 detectors in the 95GHz band, almost 10x the 288 detectors of a Keck style 95GHz receiver. Bicep3 serves as our prototype instrument leading to the eventual

replacement of the Keck array with Bicep array.

The Bicep array is a funded experiment which will replace the Keck Array for multifrequency observations. Using the more powerful Bicep3 style receivers, Bicep array will field three receivers centered at 35GHz, 95GHz, and 150GHz along with a dual band 220/270GHz receiver. The new 35GHz receiver will heavily constrain galactic synchrotron radiation past the upper limits set by WMAP's 23GHz band while the increased sensitivity at higher frequencies will allow for better constraints on dust emissions at frequencies closer to our other bands than the Planck 353GHz data.

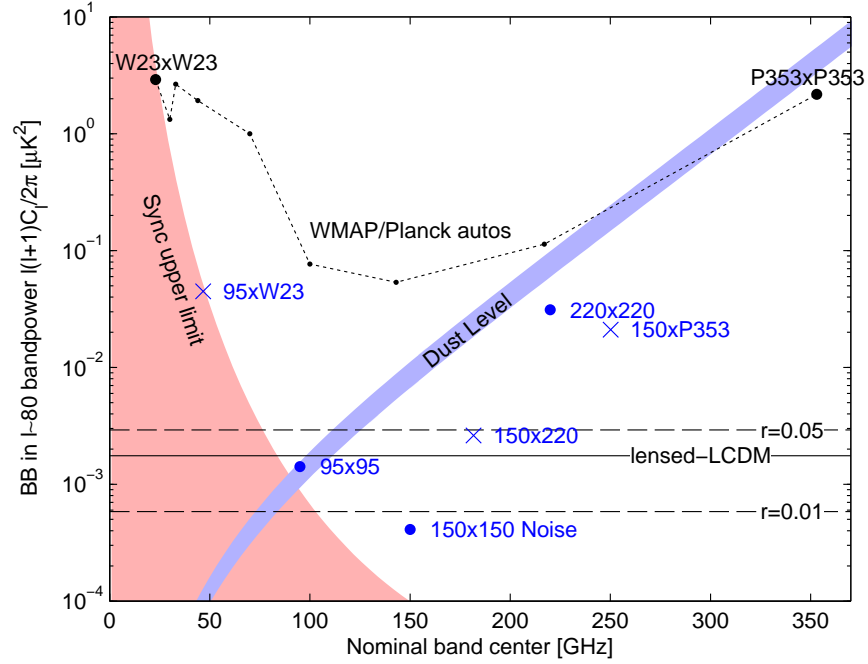
## 4 Multifrequency Observations

CMB polarization experiments must be able to separate polarized foreground signals from those imprinted on the CMB. Although these signals can be minimized by selection of observing area their emissions must be constrained and accounted for. As shown by the 2014 joint analysis between Bicep2/Keck and Planck, constraints on these models have significant impact on the interpretation of any observed excess signal. By expanding observations into multiple frequencies, the Keck array has further constrained these dust emissions as well as emissions due to galactic synchrotron as shown in Figure 4.

### 4.1 Polarized Dust

Polarized emissions from galactic dust provide an excess BB signal on top of that expected from  $\Lambda$ CDM and gravitational lensing which (although low in power) are significant compared to the IGW signal. These emissions depend significantly on frequency, exhibiting a power law like dependence. As shown in Planck XXII [14] the spectral energy distribution of galactic dust can be described by a modified blackbody spectrum





**Figure 4:** Noise levels in the Bicep/Keck combined observations through the 2015 observing season. The shaded regions show constraints on the signal level of galactic foregrounds and signal levels corresponding to select values of  $r$  are shown at the bottom. All values are shown in the  $\ell = 80$  bandpower where the IGW signal is expected to be maximal.

$$I_d = A_d \nu^\beta B_\nu(T_d) \quad (3)$$

Where  $A_d$  is an amplitude at some frequency and  $\beta_d > 0$  is the spectral index of dust emission and  $B(T)$  is the standard blackbody spectrum. In order to fully constrain the dust signal we must complement this intensity spectrum with a description of the dust's spatial behavior

$$D_\ell \propto \ell^\alpha \quad (4)$$

where  $D_\ell = C_\ell \frac{\ell(\ell+1)}{2\pi}$ . The parameters in these equations model the dust contribution to polarization signal in our field. As Equation 3 shows this signal is brighter at higher frequencies. We therefore use the Planck 353GHz maps to set these dust parameters and extrapolate to our observed frequencies. This necessarily means that any uncertainty contained in the high frequency observations is magnified due to the power law behavior.

Figure 4 shows the noise uncertainty and signal levels the  $\ell = 80$  bandpower where the IGW signal is expected to peak. The low 150x150 noise allows us to detect excess signal with high significance. However, the high P353xP353 noise as compared to dust signal does not provide significant enough constraining power to separate dust signal from potential IGW signal in the 150GHz band. The 220x220 point shows preliminary numbers from our 2015 observing season in which the Keck array operated with two 220GHz receivers and provides similar constraining power to the Planck 353GHz data while being closer to our main observing bands. Two additional 220GHz receivers were added for the 2016 observing season and observations at 270GHz will begin in the current 2017 season. These observations will allow us to produce continually improving constraints on dust in our field.

## 4.2 Galactic Synchrotron

An upper limit for an additional foreground signal is shown in Figure 4. Rather than increasing in intensity at higher frequencies, polarized emissions from galactic synchrotron radiation are strongest at low frequencies. We model synchrotron emission intensity as

$$I_\nu = A_s \nu^{\beta_s} \quad (5)$$

where  $\beta_s < 0$  describes the fall off in intensity with frequency and  $A_s$  is the amplitude. The angular power spectrum of galactic synchrotron follows the same form as dust (Equation 4). The points shown in Figure 4 mark the upper limit of synchrotron emissions as the noise level of current observations in these bands is not sufficient for detection. Models of the contribution due to synchrotron do not predict significant contamination at frequencies upwards of 150GHz due to the strong frequency dependence.

## 5 Bicep Array

### 5.1 Experiment Overview

Bicep Array will deploy four Bicep3 style receivers in four frequency bands with first light expected in 2019. Building on the success of previous experiments, Bicep Array will continue to use the antenna coupled transition edge sensor (TES) bolometers developed by JPL. These detectors leverage the strong temperature-dependence of superconductors near their transition point to detect extremely faint signals. The signal from these bolometers is then amplified by a series array of superconducting quantum interference devices (SQUIDs). We use cold optics to minimize the out of band extraneous signal incident on our detectors.

## 5.2 Cryogenic Considerations

The cryogenic operating temperatures of our detector and optics system require a specially designed housing to thermally isolate the focal plane as much as possible. As with Keck and Bicep3 the Bicep Array cryostats will consist of a nested ambient temperature vacuum jacket, a 50K radiation shield, and a 4K optics tube. This nested approach protects the low temperature optics from absorbing significant radiative power since  $P \propto (T_H^4 - T_C^4)$ . In addition to radiation shielding, there is an additional trade-off between structural support and heat conduction between the temperature shells. A low conductive loading may introduce significant deviation in pointing direction while more structural supports increase the thermal load on the cooling system.

## 5.3 Thermal Loading

The cryostat is operated at high vacuum ( $\approx 10^{-5}$  mTorr) which largely eliminates conductive loading due to residual gasses. The two largest contributions are due to conduction across the structural supports and radiation from higher temperature stages. Of these, the radiation absorbed by the lower temperature shells has the highest possible contribution. Radiative heat transfer for cryogenic systems is a well studied subject. We follow Ekin's textbook on cryogenic engineering to estimate this thermal load according to:

$$P = \sigma A_c E (T_h^4 - T_c^4) \quad (6)$$

where the  $h$  and  $c$  subscripts denote the hot and cold surfaces respectively, and  $E$  depends on the emissivity of the surfaces as

$$E = \frac{1}{1/\epsilon_h + 1/\epsilon_c - 1} \quad (7)$$

|                   |                          |
|-------------------|--------------------------|
| Rough Aluminum    | $\epsilon = 0.07 - 0.1$  |
| Polished Aluminum | $\epsilon = 0.04 - 0.06$ |
| Aluminized Mylar  | $\epsilon = 0.035$       |
| OFHC Copper       | $\epsilon = 0.65$        |

**Table 2:** Typical emissivity values for materials used in the construction of the Bicep Array cryostats.

Equation 6 is valid for parallel plates, and concentric cylinders assuming specular reflection regime. Table 2 shows typical emissivities for the materials used in the Bicep Array cryostats. In addition to the radiative load, the structural supports conduct heat between the temperature shells according to the Fourier heat transfer law:

$$\dot{Q} = \int_{T_1}^{T_2} \frac{A}{L} K(T) dT \quad (8)$$

The large temperature change across the support struts require a detailed knowledge of how the materials' thermal conductivity changes with temperature. NIST has a collection of thermal conductivity measurements for commonly used materials in cryogenic design including Aluminum and Titanium. Bicep Array also uses carbon fiber struts as structural supports, however this material is non-standardized which makes determining the thermal properties difficult, especially at high temperatures where few conductivity measurements exist. We therefore supplement published low temperature measurements with direct measurements of conducted power in a test cryostat system.

In estimating the thermal load, our ultimate objective is to ensure that the innermost 4K shell is able to sufficiently cool the second stage He3 fridge which is responsible for cooling the TES detectors down to transition. The 50K stage serves as an intermediate shield that prevents radiation from the 300K vacuum jacket from directly illuminating any 4K element. However, intercepting this radiation warms the 50K stage which is also responsible for thermally sinking an alumina filter. We therefore seek to minimize the thermal load on both the 50K and 4K stages.

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