

# A Novel Method for Removing the Cyclic Frequency Shift due to the Earth's Motion about the Sun from Astronomical Continuous-Wave Data

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By building on the method established in [5], this paper primarily seeks to introduce the Gustafson algorithm for removing frequency/phase modulation from a signal. The Gustafson algorithm uses the time series data as its input and outputs the original single tone carrier frequency. This method is powerful because it works in the frequency rather than time domain — which easily allows for filtering of the signal. Here, we will focus on the Gustafson algorithm's applications to LIGO and the search for gravitational waves; nevertheless, there are many other possible applications for the algorithm such as radio wave astronomy and communications systems using frequency modulated signals. Since one of the pertinent issues in any detection algorithm is signal adulteration by noise, we shall also introduce a novel approach to deal with correlated noise from a multichannel perspective.

## DIVISION OF LABOR

**Curtis Rau** Curtis was the most familiar with the topic and therefore performed the role of group leader identifying topics and problems of interest to be researched. He was also in charge of writing the algorithm section and the following discussion of the results. The c++ implementation of the algorithm was also done by him. He presented his part in the class presentation.

- organized topics and problems of interest, directed group research
- wrote algorithm and discussion of results
- made figures for presentation and paper
- wrote rough draft and template for paper
- Wrote the C++ implementation

**Colin Gordon** Colin contributed greatly in his research and solving of the phase moulation problem for astronomical signals cause by rotation of the earth using circular orbit approximation. He wrote the introduction to the paper, advised Sam on his section for noise and error analysis, and helped review and edit the paper as a whole. He presented his work in class.

- researched and solved phase modulation of astronomical signal caused by rotation of earth using circular orbit approximation
- wrote parts of the introduction
- reviewed paper and edited

**David Witalka** David researched and wrote the sections concerning the Nyquist frequency and Carson's rule. He also added the section pertaining to the speed of the Gustafson Algorithm. David edited and reworded the introduction and helped review and edit the paper as a whole. He also created and organized the slides presented in class.

- created and organized presentation slides
- researched and edited sections on sampling theorem, Nyquist Frequency and aliasing, Carson's rule
- reviewed paper and edited

**Samuel Akwei-Sekyere** Same researched and wrote the section on noise and error analysis, and wrote the results. He helped greatly in rewording problem sections through out the paper, and greatly expanded the discussion of all topics. Sam played a major role in reviewing and editing others' sections. He presented his work in class.

- reviewed and expanded paper (all sections)
- researched and wrote sections on error analysis, wrote the results-discussion and conclusion, wrote parts of the introduction, MATLAB code for noise section

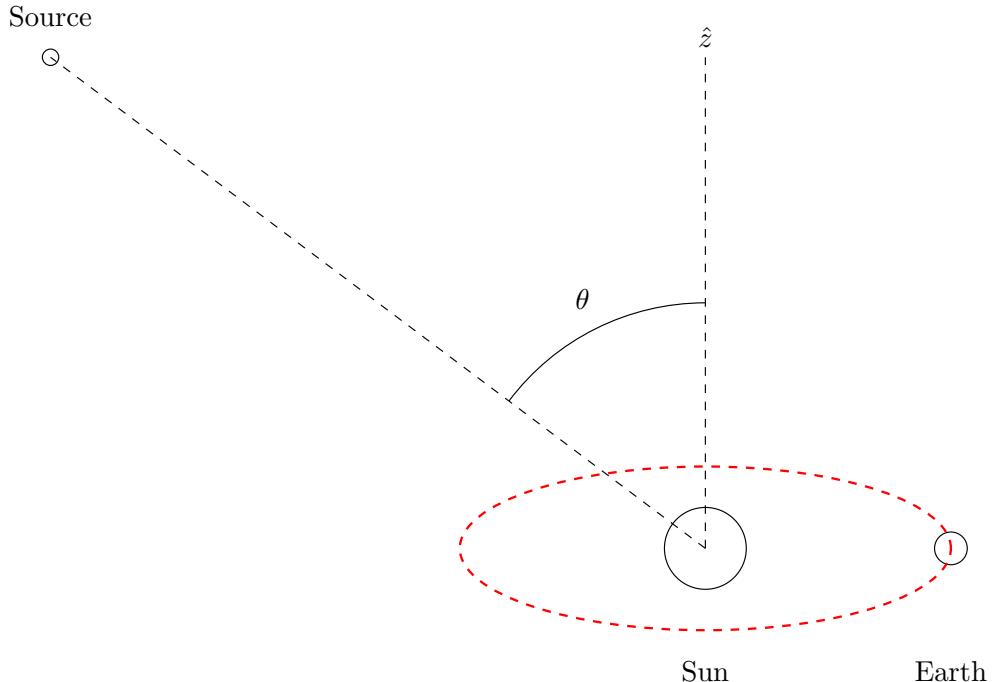
## I. INTRODUCTION

General Relativity postulates that information about the spatial distribution of matter in the universe is transmitted via a gravitational field [4]. Information regarding changes in mass distribution of the universe propagate through space-time away from the location of that change at the speed of light. Special events, where the quadrupole moment of the mass distribution changes, emit gravitational radiation [11]. Extremely violent events in the universe can release enough gravitational radiation to be detectable from Earth.

Recently, the Laser Interferometer Gravitational-Wave Observatory (LIGO) obtained the first empirical evidence of gravitational waves. The discovered ripples in space-time by LIGO was a function of the inspiral of two black holes that eventually merged — an event that consumed approximately three solar masses [6]. It is believed that gravitational wave sources are common in the universe and sources of continuous monochromatic (single frequency) gravitational radiation are predicted to exist — at least sources with frequencies that persist over many periods [9]. Neutron stars have also been implicated as one of the prime sources of gravitational waves. These stars are expected to emit single frequency waves over long periods of time. Additionally, it has also been proposed that an appreciable fraction of detectable gravitational wave emitters are binary neutron star systems — two neutron stars orbiting around each other.

A fundamental challenge to efficient detection of gravitational waves is that orbiting binary neutron star systems invoke frequency modulations which reduce the power of the carrier frequency and distributes it to higher and lower harmonics. In essence, the power of the signal is dispersed over a range of side-band frequencies; this phenomenon is due to the Doppler effect arising from the Earth's motion relative to the source of the wave. Since the noise observed in the LIGO data is approximately Gaussian and of low signal-to-noise ratio, the extraction of reliable information about gravitational waves is arduous.

In order to calculate the carrier frequency of the wave, consider the wave as one emitted from a source that is megaparsecs away (potentially in other galaxies). By virtue of its distance, the spherical coordinates will not change significantly and the source will move relative to the sun with nearly constant velocity over the time we collect data. By the same token, the Sun will experience the monochromatic wave produced by the source as truly monochromatic, while its detection on the Earth will be in the form of a frequency modulated signal.



**FIG. 1**

By assuming that the wave travels at the speed of light, we can capture it as a Lorentz invariant 4-vector,  $\mathbf{K}$  with the carrier frequency being the first entry and the second through fourth entries represent the wave number. In order to transform it in the Earth's frame we apply an instantaneous Lorentz boost on the 4-vector; the instantaneous boost takes the direction of the motion of the Earth and the velocity of the Earth in orbit into consideration. In our implementation, the Earth's orbit is captured in the  $xy$  plane. The 4-vector is written in the perspective

that assumes the wave is a plane wave which travels towards the Sun on a trajectory that is on the  $xz$  plane:

$$\mathbf{K} = \begin{pmatrix} \gamma & -\gamma\beta & 0 & 0 \\ -\gamma\beta & \gamma & 0 & 0 \\ 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & -\sin(\omega_e t + \phi_e) & \cos(\omega_e t + \phi_e) & 0 \\ 0 & -\cos(\omega_e t + \phi_e) & -\sin(\omega_e t + \phi_e) & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} \omega_0/c \\ -k \sin \theta \\ 0 \\ -k \cos \theta \end{pmatrix} \quad (1)$$

$$= \begin{pmatrix} \gamma(\omega_0/c - \beta k \sin \theta \sin(\omega_e t + \phi_e)) \\ \gamma(-\beta \omega_0/c + k \sin \theta \sin(\omega_e t + \phi_e)) \\ \gamma k \sin \theta \cos(\omega_e t + \phi_e) \\ -k \cos \theta \end{pmatrix}, \quad (2)$$

where  $\theta$  is the polar angle from the axis perpendicular to the orbit of Earth,  $\gamma$  is the Lorentz factor from special relativity,  $\omega_e$  is the angular frequency of the Earth's rotation about the Sun,  $k$  is the magnitude of the wave number,  $c$  is the speed of light and  $\beta$  is the velocity of Earth's orbit divided by the speed of light.

From this formulation we can infer that the instantaneous frequency in the Earth's frame can be given by:

$$\omega(t) = \omega_0 \gamma (1 - \beta \sin \theta \sin(\omega_e t + \phi_e)). \quad (3)$$

Since the phase refers to the time integral of the frequency, the expression for the wave being sought after is the real part of  $h(t)$ :

$$h(t) = h_0 e^{i(\omega' t + \Gamma \cos(\omega_e t + \phi_e))} \quad (4)$$

$$\omega' = \gamma \omega_0 \quad (5)$$

$$\Gamma = \frac{\gamma \beta \omega_0 \sin \theta}{\omega_e} \quad (6)$$

where  $h_0$  is the amplitude of the wave,  $\Gamma$  is the modulation index,  $\omega_0$  is the frequency of the wave to be detected,  $\phi_e$  is the relative phase difference between the Earth's rotation about the Sun and the wave to be detected.

An important note is that  $\Gamma$  is a function of only the azimuthal angle of the source with respect to the normal of the galactic plane. This algorithm will also work perfectly well with any frequency modulated wave produced similar to the above. One likely case scenario is that when the compact stars are in a binary system, one or both objects produce waves but the stars rotate around each other causing an analogous phase modulation. Withal, there could be applications for the Gustafson algorithm in areas outside of gravitational waves in areas such as radio wave astronomy and communications using frequency modulated signals. We attempt to keep the discussion of the algorithm and the results as general as possible for this reason.

## II. THE GUSTAFSON ALGORITHM

The goal of this paper is to propose a search algorithm that can efficiently detect the presence of frequency modulated waves, while being robust against noise. A brute force approach to accomplish this is to take the inner product of

$$h(t) = h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \quad (7)$$

with the data for different values of  $\omega_0$ ,  $\omega_1$ ,  $\phi_0$ ,  $\phi_1$ , and  $\Gamma$  where  $\phi_0$  is the phase of the wave to be detected at  $t = 0$ , with  $\omega_1$  and  $\phi_1$  being generalizations  $\omega_e$  and  $\phi_e$  respectively. This is effectively a Fourier transform as a function of the four search parameters:

$$\hat{f}(\omega_0, \omega_1, \phi_1, \Gamma) = \int f(t) e^{-i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} dt. \quad (8)$$

It is expected that this approach results in a something like a delta function, or at least a Dirichlet kernel because of discretization. The amount of time it will take to perform this computation is  $\mathcal{O}(N_t N_{\omega_0} N_{\omega_1} N_{\phi_1} N_{\Gamma})$ , where  $N_t$  denotes the number of data points, and the other  $N_x$  is the number of steps in  $x$  that the search will take.

Typically for a Discrete Fourier Transform (DFT) the number of steps in frequency is equal to the number of data points:  $N_t = N_{\omega_0} = N_{\omega_1}$  [12]. Physically, this infers that the frequency resolution of the DFT is proportional to the duration of the time series over which the DFT is taken. Thus the computation time for the brute force method goes as  $\mathcal{O}(N_t^3 N_{\phi_1} N_{\Gamma})$ .

The case when the signal is buried deep within non-coherent noise is of particular interest. A conventional method to improve the signal to noise ratio (SNR) is by increasing the number of data points; this assumes the signal is more periodic than noise. The problem with taking more data however is that computational time of the brute force method is proportional to  $N_t^3$ . We will now introduce the Gustafson Algorithm which is roughly a factor of  $N_t$  faster than the brute force method.

To derive the Gustafson Algorithm we start by assuming the data  $f(t)$  takes on the form of a complex frequency modulated wave. Although the recorded data is real, it will be shown later that starting with the complex wave is perfectly valid. First, we demodulate the wave and take the Fourier transform of both sides:

$$h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} = f(t) \quad (9)$$

$$h_0 e^{i\phi_0} e^{i\omega_0 t} = f(t) e^{-i\Gamma \cos(\omega_1 t + \phi_1)} \quad (10)$$

$$h_0 e^{i\phi_0} \mathcal{F}_t [e^{i\omega_0 t}] = \mathcal{F}_t [f(t) e^{-i\Gamma \cos(\omega_1 t + \phi_1)}]. \quad (11)$$

Define  $\mathcal{F}_t[f(t)](\omega) = \hat{f}(\omega)$ . By invoking the convolution theorem we yield,

$$\frac{h_0 e^{i\phi_0}}{\sqrt{2\pi}} \delta(\omega_c - \omega) = \hat{f}(\omega) * \mathcal{F}_t [e^{-i\Gamma \cos(\Omega t + \phi)}]. \quad (12)$$

The Jacobi-Anger Expansion, which comes up often in this paper, has the form:

$$e^{iz \cos(\theta)} = \sum_{n=-\infty}^{\infty} e^{in\pi/2} J_n(z) e^{in\theta}. \quad (13)$$

By invoking the previously mentioned expansion, equation 12 becomes

$$\frac{h_0 e^{i\phi_0}}{\sqrt{2\pi}} \delta(\omega - \omega_0) = \hat{f}(\omega) * \mathcal{F}_t \left[ \sum_{n=-\infty}^{\infty} e^{in\pi/2} J_n(-\Gamma) e^{in(\omega_1 t + \phi_1)} \right]. \quad (14)$$

Even order Bessel functions are even, and odd order Bessel functions are odd. Concretely,

$$J_n(-z) = (-1)^n J_n(z). \quad (15)$$

By employing this fact we can observe that

$$\mathcal{F}_t \left[ \sum_{n=-\infty}^{\infty} e^{in\pi/2} J_n(-\Gamma) e^{in(\omega_1 t + \phi_1)} \right] = \sum_{n=-\infty}^{\infty} (-1)^n e^{in\pi/2} e^{in\phi_1} J_n(\Gamma) \mathcal{F}_t [e^{in\omega_1 t}] \quad (16)$$

$$= \frac{1}{\sqrt{2\pi}} \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \pi/2)} J_n(\Gamma) \delta(\omega - n\omega_1). \quad (17)$$

In the previous step we used the linearity of the Fourier transform  $\mathcal{F}_t[c \cdot f(t)] = c \cdot \mathcal{F}_t[f]$ . Note that 14 becomes

$$h_0 e^{i\phi_0} \delta(\omega - \omega_0) = \hat{f}(\omega) * \left[ \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \pi/2)} J_n(\Gamma) \delta(\omega - n\omega_1) \right]. \quad (18)$$

By performing the convolution,

$$\hat{f}(\omega) \star \delta(\omega - n\omega_1) = \int \hat{f}(\omega - \tilde{\omega}) \delta(\tilde{\omega} - n\omega_1) d\tilde{\omega} \quad (19)$$

$$= \hat{f}(\omega - n\omega_1) \quad (20)$$

brings us to

$$h_0 e^{i\phi_0} \delta(\omega - \omega_0) = \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \pi/2)} J_n(\Gamma) \hat{f}(\omega - n\omega_1). \quad (21)$$

For clarity and later convenience let us have  $n \rightarrow -n$ . Let us employ a property of the Bessel functions:

$$J_{-n}(z) = (-1)^n J_n(z). \quad (22)$$

This results in the Gustafson Algorithm:

$$h_0 e^{i\phi_0} \delta(\omega - \omega_0) = 2 \sum_{n=-\infty}^{\infty} e^{-in(\phi_1 + \pi/2)} J_n(\Gamma) \hat{f}(\omega + n\omega_1) \quad (23)$$

An erroneous factor of two has been added to the algorithm which, as we will see, makes it suitable for use with real data. It is not obvious whether or not this algorithm in its current form can be used on real data because it assumed a complex waveform. This assumption allowed for a clean demodulation, which is marred if one takes the real of this function. This is due to the nonlinearity of the real operator as demonstrated by

$$\Re\{a \cdot b\} \neq \Re\{a\} \cdot \Re\{b\}; \quad a, b \in \mathbb{C}. \quad (24)$$

We can use equation (23) as a guess as to find the form that will work with a real valued waveform. Upon comparison of the Fourier transforms of the complex and real waveforms (see appendices A and B) we see they are very similar. The real looks like the complex with the addition of terms that are mirrored about  $\omega = 0$ .

What would we have to do to get the Fourier transform of the real to look exactly like the Fourier transform of the complex wave? Unsurprisingly, not much because they are such similar functions. All we would have to do is drop the extra mirrored terms.

We would only have to multiply each term by some phase factor to get the real Fourier transform to look like the sum of two complex Fourier transforms. After adding in the phase factors we plug the Fourier transform of the real wave into the Gustafson Algorithm which will yield two delta functions instead of one (because the Fourier transform of the real is the sum of two complex wave Fourier transforms). One delta function will be at  $+\omega_0$ , this is the one of interest, and the other will be at  $-\omega_0$ . If we carry the search out over only positive values of  $\omega_0$  then we can throw out all terms in red of the Fourier transform of the real wave because they are inconsequential to the algorithm. Finally, the phase terms that were added to the Fourier transform of the real to take on the morphology of the complex can be incorporated into the Gustafson Algorithm.

### A. Discussion

It is not explicit how The Gustafson Algorithm, Eq. 23, can be used as a search algorithm capable of determining the values  $h_0, \omega_0, \omega_1, \phi_0, \phi_1, \Gamma$ . In this section we seek to make obvious what Eq. 23 can do.

Assume the data takes on the real form of a frequency modulated wave.

$$f(t) = \Re \left\{ h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \right\} \quad (25)$$

Initially we do not know the values of  $\omega_0, \omega_1, \phi_0, \phi_1, \Gamma$ . To determine these values, we will try the values  $\tilde{\omega}_0, \tilde{\omega}_1, \tilde{\phi}_1, \tilde{\Gamma}$  in the Gustafson Algorithm.

$$G_{\tilde{\omega}_0, \tilde{\omega}_1, \tilde{\phi}_1, \tilde{\Gamma}} [\hat{f}(\omega)] = \sum_{n=-\infty}^{\infty} e^{-in(\tilde{\phi}_1 + \pi/2)} J_n(\tilde{\Gamma}) \hat{f}(\omega + n\tilde{\omega}_1) \quad (26)$$

To see what this will yield we substitute the Fourier transform of the real of the waveform, equation (B6), into the above equation

$$\begin{aligned} G_{\omega_0, \omega_1, \phi_1, \Gamma} [\hat{f}(\omega)] &= \frac{h_0}{2} \sum_{m,n=-\infty}^{\infty} e^{i(m\phi_1 - n\tilde{\phi}_1)} J_n(\tilde{\Gamma}) J_m(\Gamma) \\ &\times [e^{i(\phi_0 + (m-n)\pi/2)} \delta(\omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1) + e^{-i(\phi_0 + (m+n)\pi/2)} \delta(\omega + \omega_0 + n\tilde{\omega}_1 - m\omega_1)] \end{aligned} \quad (27)$$

We shall treat this in cases.

**Case 1:** Let us consider the most likely case when  $\omega \neq \omega_0$  and  $\omega_1 \neq \tilde{\omega}_1$ . Let us examine the term  $\delta(\omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1)$  in the above equation. This delta function appears in the double sum and is a function of both  $m$  and  $n$ , so the  $n\tilde{\omega}_1 - m\omega_1$  is changing over the sums whereas the value  $\omega - \omega_0$  is constant (not a function of  $m$  or  $n$ ). This delta is non-zero only if

$$0 = \omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1 \quad (28)$$

$$\omega_0 - \omega = n\tilde{\omega}_1 - m\omega_1 \quad (29)$$

This can happen twice or never. By the Euclidean Algorithm, the delta will only be non-zero if

$$\gcd(\tilde{\omega}_1, \omega_1) = \omega_0 - \omega \quad (30)$$

If this is false, then the sum is zero. If it is true then call the values for  $m$  and  $n$  that make it true  $\pm m_o$  and  $\mp n_o$ . So then the sum is

Placeholder for equation

**Case 2:** Now consider the case when  $\omega = \omega_0$  and  $\omega_1 \neq \tilde{\omega}_1$ . The term  $\delta(\omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1)$  will only be non-zero if

$$\tilde{\omega}_1 = \frac{m}{n} \omega_1 \quad (31)$$

Again, this can happen twice or never. If this is false, then the sum is zero.

Placeholder for equation

**Case 3:** Now consider the case when  $\omega \neq \omega_0$  and  $\omega_1 = \tilde{\omega}_1$ . The term  $\delta(\omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1)$  will only be non-zero if

$$\omega - \omega_0 = (n - m)\tilde{\omega}_1 \quad (32)$$

This too can happen twice or never. If this is false then the sum is zero.

**Cases 1-3:** In all of these cases

**Case 4:** Now consider the case when  $\omega = \omega_0$  and  $\omega_1 = \tilde{\omega}_1$ . The term  $\delta(\omega - \omega_0 + n\tilde{\omega}_1 - m\omega_1)$  will only be non-zero if  $m = n$  which will happen an infinite number of times. Thus, the sum becomes

$$\frac{h_0}{2} e^{i\phi_0} \delta(0) \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \tilde{\phi}_1)} J_n(\tilde{\Gamma}) J_n(\Gamma) \quad (33)$$

**Case 4a:** Now if we let  $\Gamma = \tilde{\Gamma}$  and  $\phi_1 = \tilde{\phi}_1$  then the sum becomes

$$\frac{h_0}{2} e^{i\phi_0} \delta(0) \sum_{n=-\infty}^{\infty} J_n^2(\Gamma) = \frac{h_0}{2} e^{i\phi_0} \delta(0) \quad (34)$$

So if we guess every search parameter correctly,  $\omega = \omega_0$ ,  $\omega_1 = \tilde{\omega}_1$ ,  $\Gamma = \tilde{\Gamma}$ , and  $\phi_1 = \tilde{\phi}_1$ , then we recover the

### III. IMPLEMENTATION OF THE GUSTAFSON ALGORITHM

#### A. Calculating Bessel Functions

Known problems with the algorithm include calculating large order Bessel functions. The Bessel functions of natural number order are given by

$$J_n(z) = \sum_{m=0}^{\infty} \frac{(-1)^n}{m!(m+n)!} \left(\frac{z}{2}\right)^{2m+n}. \quad (35)$$

Hence, calculating the  $n$ -th Bessel function involves calculating factorials greater than  $n!$ . The highest  $n$  for which we can store  $n!$  in an unsigned long integer type is 20:

$$2^{64} - 1 = 18,446,744,073,709,551,615 \quad (36)$$

$$< 21! = 51,090,942,171,709,440,000. \quad (37)$$

If we are willing to sacrifice some precision to truncation, we can use a double precision floating point integer which allows values upto  $1.797,693,134,862,315,7 \cdot 10^{308}$

$$170! \approx 7.25 \cdot 10^{306} < 1.797,693,134,862,315,7 \cdot 10^{308} < 171! \approx 1.24 \cdot 10^{309} \quad (38)$$

This is not acceptable because  $|J_n(z)| \leq 1$ . Therefore, when performing the sum we need precision all the way down to — at the very least — 0.1, which would require a mantissa with roughly 308 digits. This is not only impractical, but it is by far one of the slowest ways one could calculate values of the Bessel functions. This can be avoided by calculating the Bessel functions using Bessel's equation 39 [7].

$$x^2 J''_n(x) + x J'_n(x) + (x^2 - n^2) J_n(x) = 0 \quad (39)$$

The standard prescription for numerically solving differential equations is first to discretize the independent variable  $x \rightarrow x_i = x_{min} + i \cdot \delta x$ , and then employ the limit definitions of the derivatives. We will use the three point definitions because their error goes as  $O(\delta x^2)$  rather than using the two point definition with an error that goes as  $O(\delta x)$  [? ]. Further, because there are only one order of Bessel functions in Bessel's equation, the order is implied by the  $n$  that appears, so we will drop the subscript of  $n$ ,  $J_n \rightarrow J$ , and adopt the notation  $J(x_i) = J_i$ .

$$J'_i = \frac{J_{i+1} - J_{i-1}}{2\delta x} \quad (40)$$

$$J''_i = \frac{J_{i-1} - 2J_i + J_{i+1}}{\delta x^2}. \quad (41)$$

Therefore the discretized form of Bessel's equation is

$$x_i^2 \frac{J_{i-1} - 2J_i + J_{i+1}}{\delta x^2} + x \frac{J_{i+1} - J_{i-1}}{2\delta x} + (x^2 - n^2) J_i = 0 \quad (42)$$

$$\left( \frac{x_i^2}{\delta x^2} - \frac{x_i}{2\delta x} \right) J_{i-1} + \left( \left( 1 - \frac{2}{\delta x^2} \right) x_i^2 - n^2 \right) J_i + \left( \frac{x_i^2}{\delta x^2} + \frac{x_i}{2\delta x} \right) J_{i+1} = 0 \quad (43)$$

where we have collected like terms of  $J_i$ .

For convenience we define coefficients to simplify the above equation.

$$a_i = \frac{x_i^2}{\delta x^2} - \frac{x_i}{2\delta x} \quad (44)$$

$$b_i = \left( 1 - \frac{2}{\delta x^2} \right) x_i^2 - n^2 \quad (45)$$

$$c_i = \frac{x_i^2}{\delta x^2} + \frac{x_i}{2\delta x} \quad (46)$$

These coefficients can be simplified by using the definition for  $x_i$ , and taking  $x_{min} = 0$ .

$$a_i = i \left( i - \frac{1}{2} \right) \quad (47)$$

$$b_i = i^2 (\delta x^2 - 2) - n^2 \quad (48)$$

$$c_i = i \left( i + \frac{1}{2} \right) \quad (49)$$

In effect, the discretized Bessel's equation with simplified coefficients is

$$a_i J_{i-1} + b_i J_i + c_i J_{i+1} = 0 \quad (50)$$

By writing out the system of equations explicitly, we have:

$$b_1 J_1 + c_1 J_2 = -a_1 J_0 \quad (51)$$

$$a_2 J_1 + b_2 J_2 + c_2 J_3 = 0 \quad (52)$$

$$a_3 J_2 + b_3 J_3 + c_3 J_4 = 0 \quad (53)$$

$$\vdots \quad (54)$$

$$a_{n-1} J_{n-2} + b_{n-1} J_{n-1} + c_{n-1} J_n = 0 \quad (55)$$

$$a_n J_{n-1} + b_n J_n = -c_n J_{n+1} \quad (56)$$

This suggests Bessel's equation can be written as a matrix equation:

$$\begin{pmatrix} b_1 & c_1 & & & \\ a_2 & b_2 & c_2 & & \\ a_3 & b_3 & c_3 & & \\ \vdots & \vdots & & & \\ a_{n-1} & b_{n-1} & c_{n-1} & & \\ & a_n & b_n & & \end{pmatrix} \begin{pmatrix} J_1 \\ J_2 \\ J_3 \\ \vdots \\ J_{n-1} \\ J_n \end{pmatrix} = \begin{pmatrix} -a_1 J_0 \\ 0 \\ 0 \\ \vdots \\ 0 \\ -c_n J_{n+1} \end{pmatrix}. \quad (57)$$

This is a tridiagonal matrix system that characterized a one dimensional differential equation. It can be easily solved using the standard Thomas Algorithm. In order to make it easily programmable, this implementation assumes all the arrays are indexed from 0. By virtue of this indexing system, we yield:

$$\begin{pmatrix} B_0 & C_0 & & & \\ A_1 & B_1 & C_1 & & \\ A_2 & B_2 & C_2 & & \\ \vdots & \vdots & & & \\ A_{N-4} & B_{N-4} & C_{N-4} & & \\ & A_{N-3} & B_{N-3} & & \end{pmatrix} \begin{pmatrix} J_1 \\ J_2 \\ J_3 \\ \vdots \\ J_{N-3} \\ J_{N-2} \end{pmatrix} = \begin{pmatrix} -A_0 J_0 \\ 0 \\ 0 \\ \vdots \\ 0 \\ -C_{N-3} J_{N-1} \end{pmatrix} \quad (58)$$

$$A_i = a_{i+1} = (i + 1) \left( i + \frac{1}{2} \right) \quad (59)$$

$$B_i = b_{i+1} = (i + 1)^2 (\delta x^2 - 2) - n^2 \quad (60)$$

$$C_i = c_{i+1} = (i + 1) \left( i + \frac{3}{2} \right) \quad (61)$$

### 1. The Thomas Algorithm

The Thomas Algorithm is a lightning fast way to solve tridiagonal systems like equation 58. The algorithm consists of two steps. First there is a forward substitution which eliminates the  $A_i$ 's, and modifies the  $B_i$ 's and the vector on the right which we will call the source vector  $S_i$ ;  $\forall i \in [0, N - 3]$ .

$$R_i = R_i - \frac{A_i}{B_{i-1}} R_{i-1}; \quad \forall i \in [1, N - 3] \quad (62)$$

$$B_i \rightarrow B_i - \frac{A_i}{B_{i-1}} C_{i-1}; \quad \forall i \in [1, N - 3] \quad (63)$$

$$S_i \rightarrow S_i - \frac{A_i}{B_{i-1}} S_{i-1}; \quad \forall i \in [1, N - 3] \quad (64)$$

Then there is a backwards substitution eliminating the  $C_i$ 's, again modifying the  $S_i$ 's. Finally we solve for  $J_i$ .

$$R_i = R_i - \frac{C_i}{B_{i+1}} R_{i+1}; \quad \forall i \in [N - 4, 0] \quad (65)$$

$$S_i \rightarrow S_i - \frac{C_i}{B_{i+1}} S_{i+1}; \quad \forall i \in [N - 4, 0] \quad (66)$$

$$J_i = \frac{S_{i-1}}{B_{i-1}}; \quad \forall i \in [N - 2, 1] \quad (67)$$

### 2. Boundary Conditions

To solve this second-order differential equation we will need two boundary conditions. The boundary condition at  $x = 0$  is simple:

$$J_n(0) = \begin{cases} 1 & : n = 0 \\ 0 & : n \neq 0 \end{cases}$$

For the other boundary condition we *could* use the asymptotic form of the Bessel functions. Theorem 5.1 in [7] states:

*For each  $n \in N$  there is a constant  $C_n \in R$  such that, if  $x \geq 1$ , then*

$$\left| J_n(x) - \sqrt{\frac{2}{\pi x}} \cos \left( x - \frac{\pi}{4}(2n + 1) \right) \right| \leq \frac{C_n}{x^{3/2}}. \quad (68)$$

This is an asymptotic expansion, so it is only valid for large values of  $x$ . It turns out the accuracy of the solution to Bessel's equation relies heavily on the accuracy of the endpoint at  $x \neq 0$ . Figures ?? and ?? illustrate this point quite well. The conclusion is that the asymptotic form of the Bessel functions is not accurate enough for even fairly large values of  $z$ . We can also do

$$J_n(z) = \frac{1}{\pi} \int_0^\pi \cos [z \sin \theta - n\theta] d\theta. \quad (69)$$

Which is computed using a Riemann sum:

$$J_n(z) \approx \frac{1}{N} \sum_{i=0}^{N+1} \cos [z \sin(i \cdot dx) - n \cdot i \cdot dx]; \quad dx = \frac{\pi}{N} \quad (70)$$

The function being integrated becomes more oscillatory as  $n$  increases. Numerical analysis shows that  $n$  gives the number of zeros the function has on the interval  $(0, \pi)$  for  $z = 0$ . For  $n = 0$  the number of zeros is given by

$\lfloor \frac{2z}{\pi} \rfloor$ . We want a way to ensure the calculation is accurate regardless of how oscillatory it is, but we also want it to be as fast as possible. So, we use these facts regarding the zeros to ensure there are roughly the same number of integration points between zeros. By taking the number of (evenly spaced) integration points to be

$$N = 100 \cdot (n + 1) \cdot \left\lfloor \frac{2z}{\pi} + 1 \right\rfloor, \quad (71)$$

we achieved excellent results (accurate consistently to 16 decimal places) for a wide range of  $z$  and  $n$ . The  $+1$ 's are to ensure accurate results when  $z$  or  $n$  are zero. The algorithm can have problems when  $J$  is extremely small. Then the function can be off by many orders of magnitude, but because it is essentially zero, it does not seem to effect our results when used as a boundary condition for solving Bessel's Differential Equation.

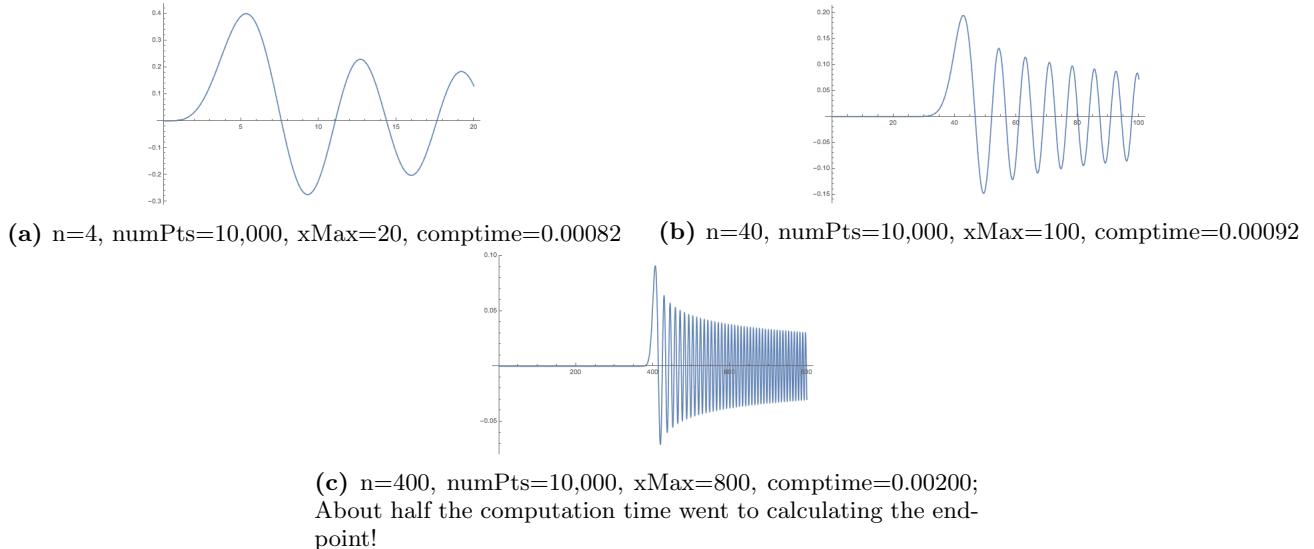
### 3. Calculating Bessel Functions: Results

The ultimate benchmark for success in calculating Bessel functions in this manner (using the Thomas Algorithm to solve Bessel's Differential Equation) is to compare it to the standard way of calculating the value at each point using the Riemann Sum (as defined in equation 70). So how much faster is it? This depends on the values of  $n$  and  $z$  because we use a dynamic number of integration points as given by eq. 71. Table I gives a typical speed increase for various parameters.

Parameters	Point by Point Using Integral	Thomas Algo. Solving Bessel's Equation	Speed Increase
$n = 3$	0.051786 s	0.00013800 s	375
$N = 500$			
$z_{max} = 20.0$			
$n = 14$	113.33 s	0.003937 s	28,785
$N = 50,000$			
$z_{max} = 100.0$			

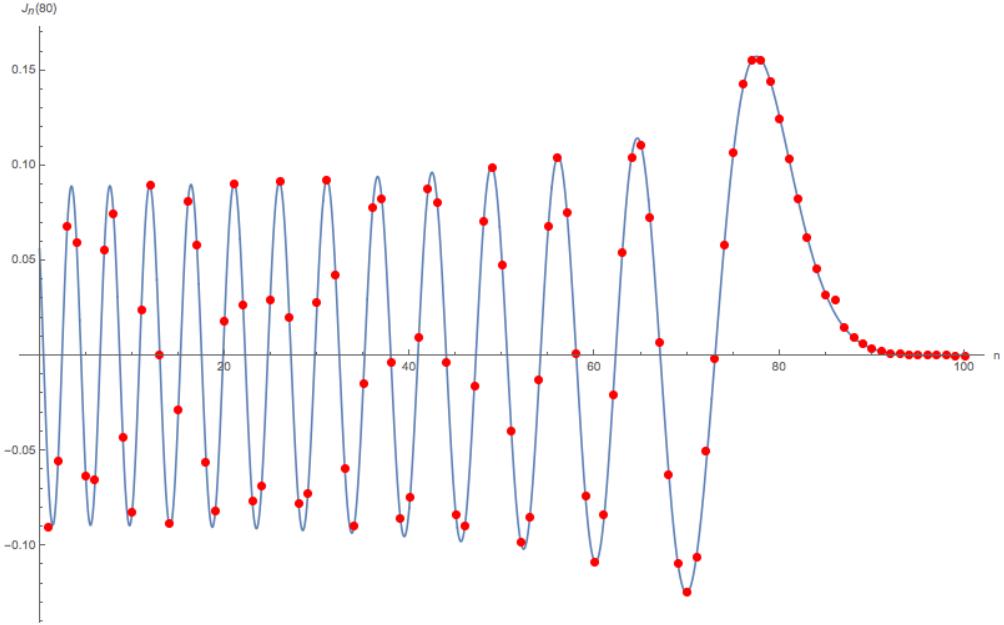
**TABLE I:**  $N$  is the number of points returned,  $n$  is the order of the Bessel function, and the points calculated were for  $z \in [0, z_{max}]$

Figure 3 provides three examples of Bessel Functions that were calculated using the Thomas Algorithm to solve Bessel's Equation.



**FIG. 3**

In Figure 4, it can be observed that the standard approach and the suggested approach are in agreement except at  $n = 0, 85$ . This is a small subset of the 400,000 data points generated in 0.33 seconds (on 5 year old a laptop). It solved the Bessel differential equation for 4000 steps in  $\Gamma \in (0, 100)$ , then saved the output, and moved on to solve Bessel's equation for the next value of  $n$  until all values of  $\Gamma$  and  $n$  were solved for.



**FIG. 4:** Plot of  $J_n(80)$  for values of  $n \in (0, 100)$ . The blue curve was generated by the BesselJ function of Mathematica. The red dots were calculated by the suggested approach.

#### IV. CARSON'S RULE AND THE NYQUIST FREQUENCY

In order too fully reconstruct the carrier frequency we would need to take an infinite sum over the sidebands. As this is impractical and unnecessary as many terms are insignificant, we need to find a way to limit the number of sidebands we are summing over. The first problem is locating the sidebands' characteristic frequency. There are infinite sideband frequencies and eventually they will be higher than the Nyquist frequency which will cause aliasing of the sideband frequency. The Nyquist frequency,  $\psi_N$ , is defined as the maximum frequency that can be accurately reconstructed and is equal to half of the sampling frequency,  $\psi_s$ .

$$\psi_N = \frac{1}{2}\psi_s \quad (72)$$

Any side band with a frequency of  $\psi_0$  higher than the Nyquist frequency, but less than the sampling frequency, will be reconstructed with a frequency of  $(\psi_s - \psi_0)$ . That is the frequency will be under-sampled and an alias frequency will be produced. When this occurs the best case scenario is that the alias frequency is unique and does not correspond to the frequency of another lower frequency sideband. In this case we can use anti-aliasing techniques to correct for the aliasing and still obtain information about the amplitude and frequency of the higher side bands. The worst case scenario is that the aliased frequency is too close to another lower frequency sideband and this will cause two problems: a systematic error will be added to the measurement of the lower side band's amplitude, and we will be unable to determine the necessary information on what the amplitude of the higher frequency would be. To minimize the effect this might have, we would like any sidebands with a frequency over the Nyquist frequency to have as low an amplitude as possible. For this we can employ Carson's rule. Carson's rule can be understood to say that almost all (roughly 98 percent) of the power for a frequency-modulated sinusoidal signal is contained within a finite bandwidth  $B_T$ , defined by:

$$B_T = 2(\Delta\psi + \psi_m) \quad (73)$$

where  $\Delta\psi$  is the peak frequency deviation of the instantaneous frequency  $\psi(t)$  from the center carrier frequency  $\psi_c$ , and  $\psi_m$  is the highest frequency in the modulating signal. In our case the highest frequency of the modulating

signal is the carrier frequency,  $\psi_c$ . From this calculation we can adjust our sampling frequency to ensure it is large enough in order to minimize aliasing. This also puts a limit on the number of sidebands we need to sum over to just those that fall within the bandwidth. The important thing to take away from this is that it takes an infinite bandwidth to perfectly transmit a phase or frequency modulated wave, regardless of how smooth it is, but in practice only a finite bandwidth is needed for an accurate approximation.

## V. ERROR ANALYSIS

While the search for monochromatic gravitational waves via the LIGO detector is relatively effortless, the adulteration of signals obtained from the detector by noise adds an appreciable number of complications to the problem. By the same token, techniques by which improved resistance of the LIGO detector to noise can be obtained has been investigated vigorously [2] [1].

### A. Noise Sources

Since the signal being searched for is relatively stable but of little amplitude, the reliability of data obtained from the detector is related to the degree of noise therein. Among others, quantum, thermal, seismic and Newtonian noise have been implicated as prime drivers of LIGO data adulteration. Since the prime noise contributor is quantum, a summation of all the noise sources are approximately Gaussian. Thus, it suffices to use the properties of Gaussian noise to understand the volatility of the data.

### B. Noise Elimination

As mentioned previously, quantum noise — in particular, thermal noise — has been implicated as a limiting factor in the sensitivity of gravitational wave interferometers and hence a prime factor in the degradation of the signal quality; this noise source is generally modeled as Gaussian [13].

By low-pass filtering or spectral subtraction, it is possible to derive signals with high SNR; however, these widely-accepted approaches require information which can be difficult to obtain. Since the frequency of the monochromatic wave being searched for is not known, low-pass filtering may result in erroneous outputs. Spectral subtraction requires a sample of the noise; nevertheless, due to the low SNR of LIGO data, it is precarious to find segments with pure noise.

In this section, we introduce an alternate approach by which Gaussian noise in data may be reduced. This method utilizes the following principle: the nature of the noise is not known but the nature of the signal is known. Our methodology consists of subtracting estimated noise from the signal via singular value decomposition of the Fourier transform of intrinsic mode functions (see appendix) and assumes the noise is correlated across different channels. A significant part of this approach is motivated by [10] and [3]. It is worth noting that because correlated Gaussian noise is prevalent, this approach may be used for a host of other signals.

#### 1. Proposed Denoising Algorithm

In an effort to reduce the noise, we shall consider data from two channels. By the assumption that the noise observed is additive (and correlated), we yield the following:

$$\xi^{(1)}(t) = f^{(1)}(t) + \epsilon^{(1)}(t) \quad (74)$$

$$\xi^{(2)}(t) = f^{(2)}(t) + \epsilon^{(2)}(t) \quad (75)$$

where  $\xi^{(n)}(t)$  is the data observed,  $f^{(n)}(t)$  is the true signal and  $\epsilon^{(n)}(t)$  is the noise obtained from the n-th channel (or gravitational wave interferometer) with  $\epsilon^{(n)}(t) \sim \mathcal{N}(0, \sigma^2)$ .

First, we disintegrate each observed signal into  $\alpha$  intrinsic mode functions via empirical mode decomposition (see appendix) and arrange them into a  $2\alpha \times T$  matrix  $\mathbf{L}$ :

$$\mathbf{L} = \begin{pmatrix} l_{1,t_1}^{(1)} & l_{1,t_2}^{(1)} & \dots & l_{1,t_T}^{(1)} \\ l_{2,t_2}^{(1)} & l_{2,t_2}^{(1)} & \dots & l_{2,t_T}^{(1)} \\ \vdots & \vdots & \vdots & \vdots \\ l_{\alpha,t_1}^{(1)} & l_{\alpha,t_2}^{(1)} & \dots & l_{\alpha,t_T}^{(1)} \\ l_{\alpha+1,t_1}^{(2)} & l_{\alpha+1,t_2}^{(2)} & \dots & l_{\alpha+1,t_T}^{(2)} \\ l_{\alpha+2,t_2}^{(2)} & l_{\alpha+2,t_2}^{(2)} & \dots & l_{\alpha+2,t_T}^{(2)} \\ \vdots & \vdots & \vdots & \vdots \\ l_{2\alpha,t_1}^{(2)} & l_{2\alpha,t_2}^{(2)} & \dots & l_{2\alpha,t_T}^{(2)} \end{pmatrix} \quad (76)$$

where  $l_{p,t_q}^{(n)}$  is the  $\left[ \left( (p-1) \bmod \alpha \right) + 1 \right]$ -th intrinsic mode function at time  $t_q$  of  $\xi^{(n)}(t)$ , and with  $T$  being the total number of data points observed.

Thereafter, the fast Fourier transform of each intrinsic mode function in  $\mathbf{L}$  is computed to obtain a matrix  $\mathcal{F}(\mathbf{L})$ . Singular value decomposition of the variance-covariance matrix obtained from  $\mathcal{F}(\mathbf{L})$  is then computed to result in the following:

$$\Sigma = \mathbf{U}\mathbf{S}\mathbf{V} \quad (77)$$

where  $\Sigma = T^{-1} [\mathcal{F}(\mathbf{L})^T - \mathbf{1}\mathbf{1}^T \mathcal{F}(\mathbf{L})^T T^{-1}]^\dagger [\mathcal{F}(\mathbf{L})^T - \mathbf{1}\mathbf{1}^T \mathcal{F}(\mathbf{L})^T T^{-1}]$  with  $(\cdot)^T$  being the transpose, the conjugate transpose  $(\cdot)^\dagger$  and  $\mathbf{1}$  a  $T \times 1$  column matrix.

The fast Fourier transforms are then decorrelated to form a matrix  $\mathbf{R}$  by employing the following:

$$\mathbf{R} = \mathbf{U}^T \Sigma \mathbf{V}. \quad (78)$$

After the decorrelation procedure above, we seek to apply a threshold  $\Delta$  which dictates the approximate percentage of statistically irrelevant information about the Fourier transform we shall discard. Prior to the thresholding procedure, the eigenvectors that contribute to approximately  $\Delta$  percent of the variance of the transform by virtue of their respective eigenvalues is estimated using the following formulation:

$$y = \operatorname{argmin}_b \left( \left| \frac{\sum_{j=1}^{(b)} \lambda_j}{\operatorname{trace}\{\mathbf{S}\}} - \frac{\Delta}{100} \right| \right) \quad (79)$$

where  $y$  is the index denoting the  $y^{\text{th}}$  eigenvalue from  $\lambda_y$  with  $\lambda = \operatorname{diag}(\mathbf{S})$ . It is essential to note that the eigenvalues are ordered from the highest to the lowest; the eigenvectors are also ordered by the position of their corresponding.

To apply the threshold, the projections with eigenvalues whose indices are less than  $y$  are retained and the rest are replaced with zero row vectors:

$$\mathbf{R}_r = \begin{cases} \mathbf{R}_r & , r < y \\ \mathbf{0} & , r \geq y \end{cases}$$

where  $\mathbf{R}_r$  the  $r^{\text{th}}$  row of the matrix  $\mathbf{R}$ . This preserves the vectors that explain closest to, but less than,  $\Delta$  percent of the total variance.

The fast Fourier transforms are then recorrelated to obtain denoised intrinsic mode functions in the following manner:

$$\tilde{\mathbf{L}} = \mathbf{V}\mathbf{R}. \quad (80)$$

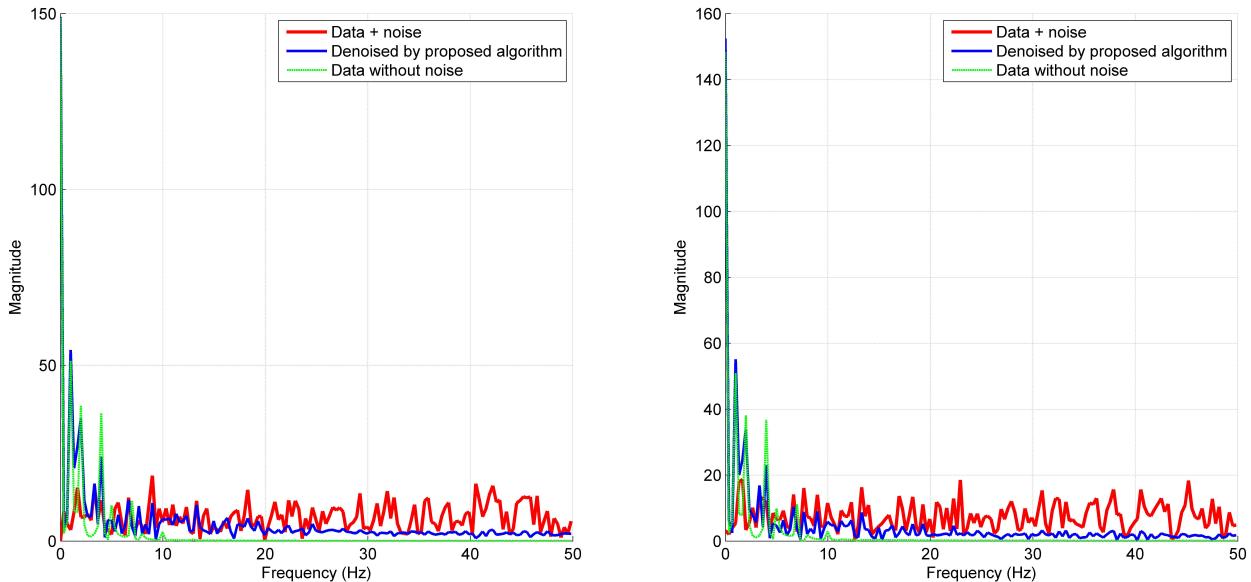
The real component of the inverse fast Fourier transform of each row of  $\tilde{\mathbf{L}}$  is then computed to obtain a matrix  $\mathbf{D}$  of similar architecture with denoised intrinsic mode functions. Thereafter, the reconstructed signals  $\hat{\xi}^{(n)}(t)$  are obtained as follows:

$$\hat{\xi}^{(n)}(t_q) = \sum_{u=(n-1)\alpha+1}^{n\alpha} \mathbf{d}_{u,q} \quad (81)$$

where  $\mathbf{d}_{u,q}$  is the value at the  $u^{\text{th}}$  row and  $q^{\text{th}}$  column of  $\mathbf{D}$ .

## 2. Discussion

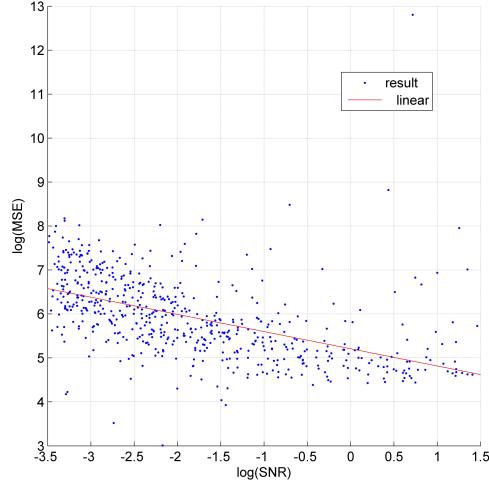
To demonstrate the potential of this procedure in removing correlated Gaussian noise, two frequency modulated signals that are phase shifted from each other are adulterated with correlated Gaussian noise. The original signals are then reconstructed using the proposed denoising algorithm with a threshold of 99 percent — that is, eigenvectors that account for closest to (but, less than) 99 percent of the variation of the data obtained from the set of multichannel intrinsic mode functions are retained. In Figure 5a, it can be easily observed that the algorithm extinguished an appreciable amount of the high frequency component that was not previously present in the data and selectively reduced the magnitude of the Fourier transform of the noise components with low frequency. Figure 5b was no different in terms of its denoising performance. It is worth noting that both results do have some discrepancy — and their discrepancies are not the same. This feature is due to the fact that, although the noise sources are correlated, the noises are different.



(a) Fourier transform of the proposed denoising approach for channel 1.  
(b) Fourier transform of the proposed denoising approach for channel 2.

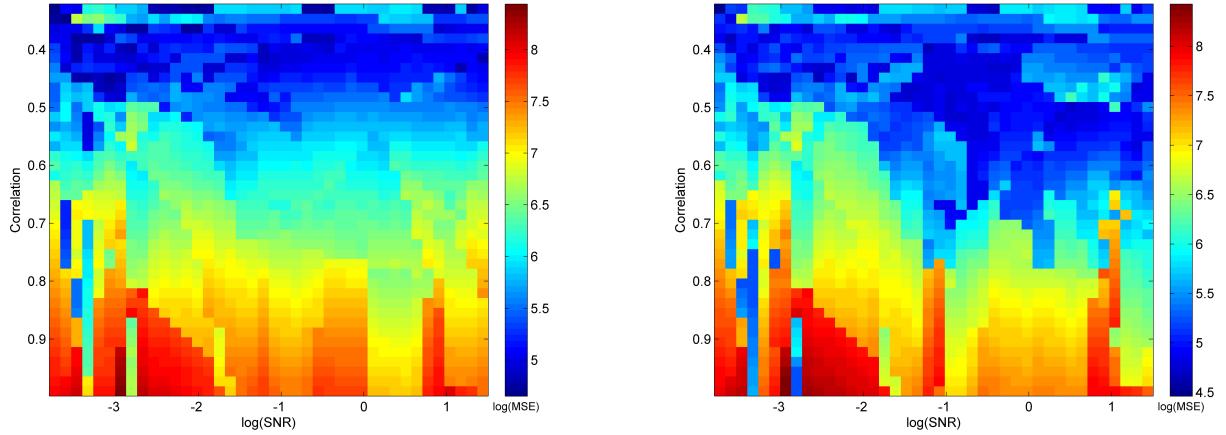
**FIG. 5:** Output of the proposed denoising approach on two channels with signals of similar frequency modulation, but shifted in time. This procedure assumes the signal being recorded is from the same source and the noise observed in each channel are correlated.

A fundamental question that needs a resolution is about how the proposed algorithm responds to different SNRs. To address this, the deviations of the Fourier transforms of reconstructed signals as a function of a myriad of SNRs were computed from the Fourier transform of the original signal. These deviation scores were measured as the mean squared errors between the reconstructed signals and the original signal. The threshold employed in this investigation was 99 percent. From Figure 6, it can be noted that the performance of the proposed algorithm increases with an increasing SNR; however, this increase in performance is limited. Since the proposed approach disregards parts of the data that are statistically irrelevant (with respect to the threshold chosen), it is plausible that parts of the original data may have been erased. This suggests that an adaptive threshold that depends on the structure of the data and the structure of the noise to be extinguished may be most appropriate for this algorithm.



**FIG. 6:** The algorithm performs better with increasing SNR; however, this performance plateaus at high SNRs. The correlation coefficient  $r^2 = -0.6543$  with  $p < 0.05$ .

Since a fundamental assumption the proposed algorithm makes is that the noise sources are correlated, it is important to investigate how the correlation affects its performance. To this end, a procedure similar to that of Figure 6 was implemented; nevertheless, a second independent variable was added — the correlation between the noise sources. The performance for denoising in each channel was not analyzed jointly. From Figure 7a, it can be noticed — as insinuated before — that the algorithm’s performance plateaus at high SNRs and progressively worsens as the same threshold is more likely to remove parts of the original data. Interestingly, the proposed approach works well with noise sources of tolerably high correlation; however, it fails to produce quality results when the correlation is really high.



(a) Response of the algorithm with changing noise correlations and SNRs for channel 1. (b) Response of the algorithm with changing noise correlations and SNRs for channel 2.

**FIG. 7:** Behavior of the proposed algorithm with respect to differing SNRs and noise correlations. This figure provides an insight to the performance of the proposed denoising algorithm (via the mean squared error) under the outlined conditions.

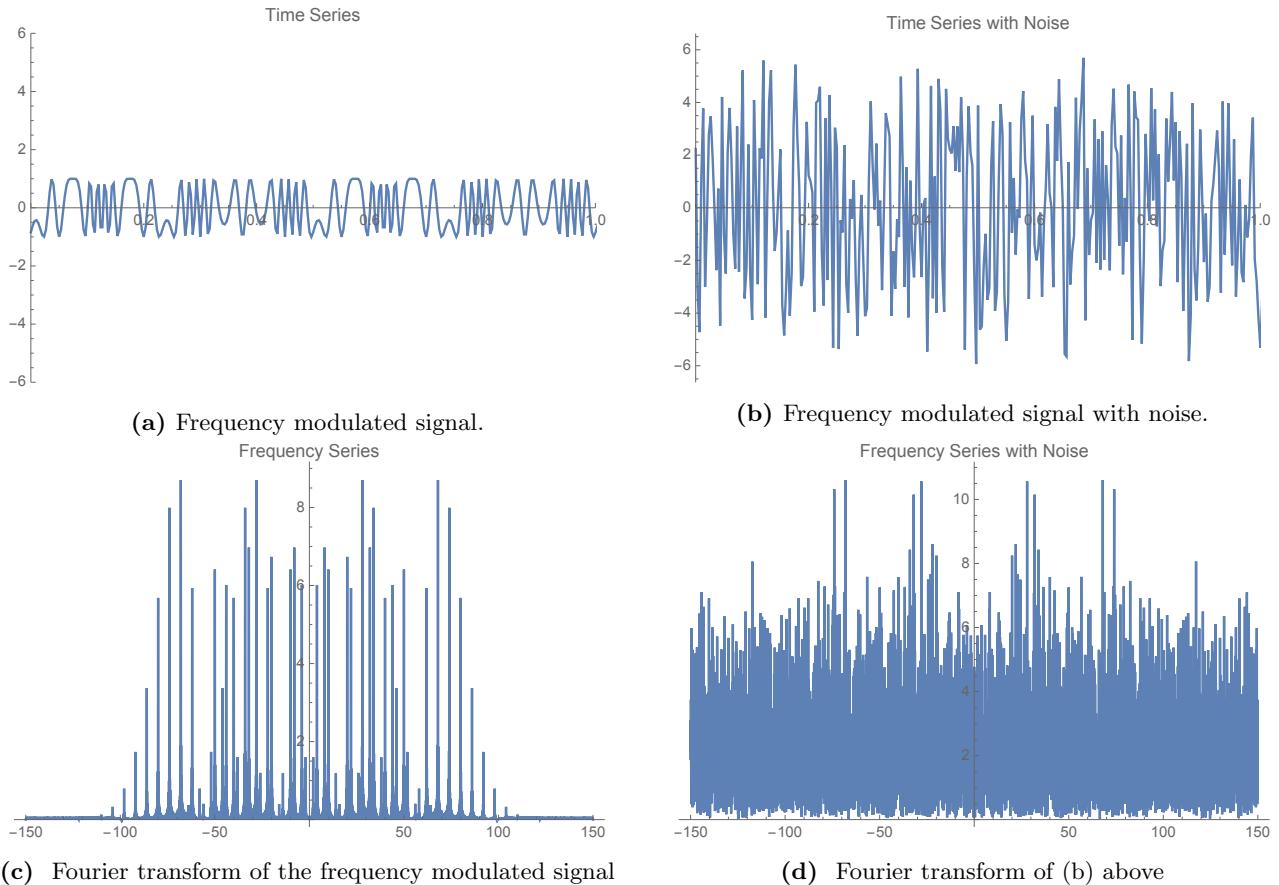
The potential of this approach has been outlined above. With a few modifications, this denoising algorithm may improve its capability and, to some extent (depending on the structure of the data and the noise), could be modified to deal with uncorrelated noise. In the spirit of demonstrating the robustness of the Gustafson algorithm against noise, the results from the search algorithm will be analyzed with no attempt to remove noise.

## VI. RESULTS AND DISCUSSION

Since frequency modulated signals may sometimes look like noise in the time domain (see appendix), it is imperative to ascertain that the Gustafson algorithm provides accurate information about the carrier frequency. For this purpose, the Gustafson algorithm was designed to be implemented in Fourier space.

By observing Figures 8a and 8b, we can see that noise added to the frequency modulated signal makes it almost impossible to visually characterize. In practice, most frequency modulated signals encountered — especially those of astronomical sources and signal susceptible to environmental noise — are of this form. A simple approach to identifying relevant features of the frequency modulated signal adulterated with noise will be to search for patterns therein. However, since the noise is sampled randomly it is nearly impossible to extract patterns without some form of transformation. The most obvious transform to employ that allows us access the frequencies that compose the noisy signal is its Fourier transform.

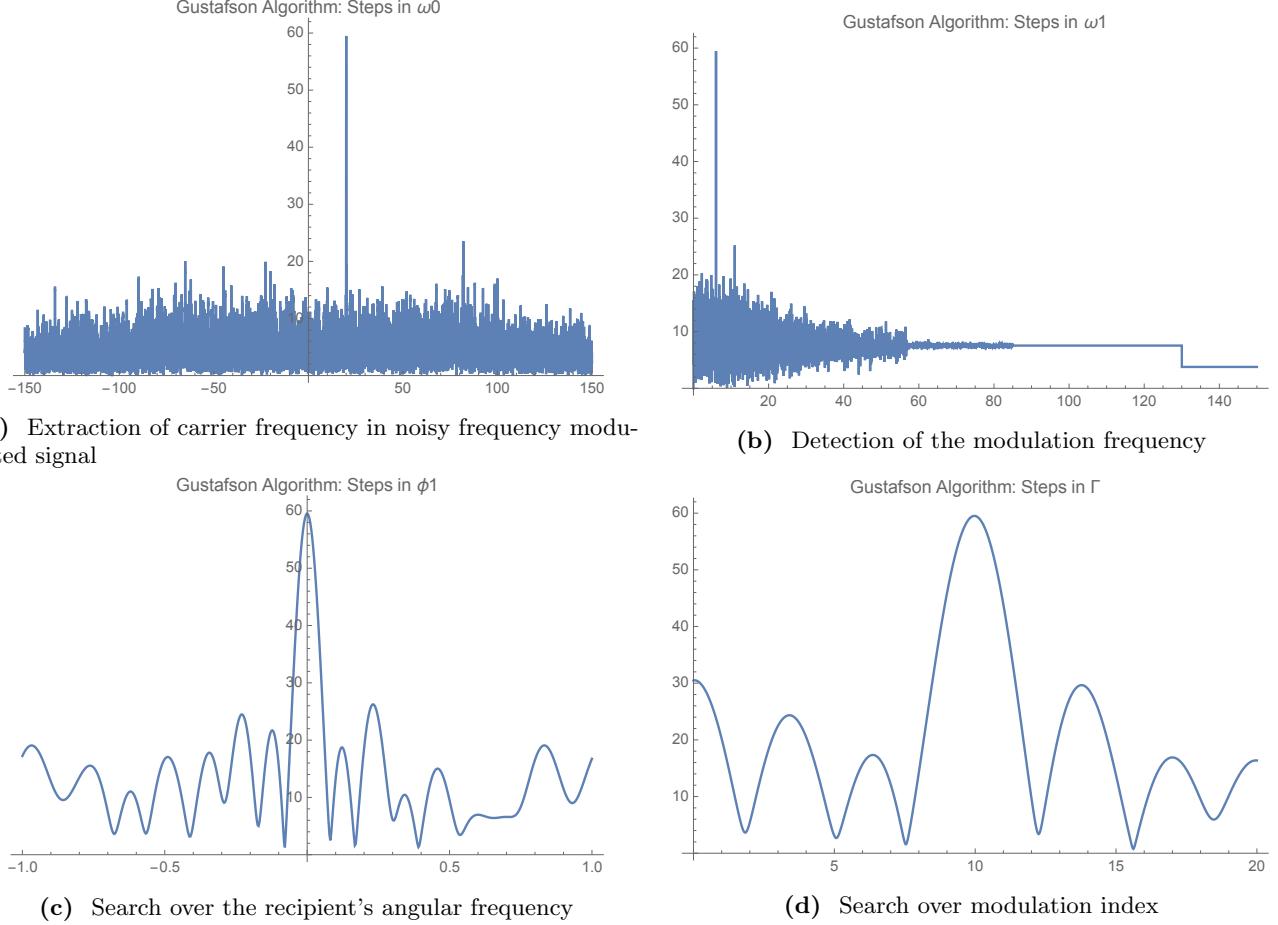
In Figures 8c and 8d the Fourier transform of the frequency modulated signal and that of the signal adulterated with noise are shown respectively. In the space of the Fourier transform, the frequencies that compose the original signal can be clearly observed. However, this task gets challenging when observing the Fourier transform of the noisy signal. Note that although a high degree of noise has been added, major peaks can still be seen. This is a function of the SNR; with a high enough noise amplitude, the Fourier transform will fail in providing useful information. This suggests that although the Fourier transform is a powerful technique when it comes to ignoring noise within a continuous signal, it is limited by the power of the noise.



**FIG. 8:** Frequency modulated signal and its noisy version in the time series and Fourier space. The sampling rate for these traces was 300Hz. The carrier frequency was 20Hz, while the modulation frequency was 6Hz. It is worth noting that these represent 10 seconds of recording with a modulation index of 10.

To test the robustness of the Gustafson algorithm against noise, we searched for the carrier frequency, modulation frequency, modulation index and the recipient's angular frequency for the noisy signal provided in Figure 8b. From Figure 9a we can see that the Gustafson algorithm appropriately separated out the signal (carrier frequency) from the noise even though it was noise obvious in the corresponding Fourier transform. It is worth noting that the search results in a transform that is asymmetric in frequency; this is a function of the fact that the Gustafson algorithm

includes a rotation in Fourier space that is opposite to that of the Fourier transform. With Figure 9b, it can be observed that the Gustafson algorithm accurately obtained modulation frequency of the noisy signal. Since the search for the carrier frequency and modulation frequency is sensitive, it is important to use small steps; this will aid in preventing one from accidentally skipping over the correct solution. Like the previous results, the Gustafson algorithm was able to extract the correct recipient's angular frequency (Fig. 9c) and the modulation index (Fig. 9d) of the frequency modulated signal. Since the search for the modulation index and recipient's angular frequency is relatively smooth, one may take tolerably larger step sizes and still converge to the right solution.



**FIG. 9:** Application of the Gustafson algorithm to the noisy frequency modulated signal illustrated in Fig. ??

## VII. CONCLUSION

In this paper, the Gustafson algorithm has been introduced as well as a novel approach to noise reduction. Although both algorithms presented have been shown to be effective, they may require substantial modifications to be useful in practice. One major draw back of the Gustafson algorithm is the selection of a range of parameters to search over. An investigation of how one can introduce learning procedures such as stochastic gradient descent might be able to provide a steady solution to this issue.

### Appendix A: The Fourier Transform of The Complex Function

$$\mathcal{F}_t \left[ h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \right] = h_0 e^{i\phi_0} \mathcal{F}_t \left[ e^{i\omega_0 t} e^{i\Gamma \cos(\omega_1 t + \phi_1)} \right] \quad (\text{A1})$$

$$= h_0 e^{i\phi_0} \mathcal{F}_t \left[ e^{i\omega_0 t} \right] \star \mathcal{F}_t \left[ e^{i\Gamma \cos(\omega_1 t + \phi_1)} \right] \quad (\text{A2})$$

$$= 2\pi h_0 e^{i\phi_0} \delta(\omega - \omega_0) \star \mathcal{F}_t \left[ \sum_{n=-\infty}^{\infty} e^{in\pi/2} J_n(\Gamma) e^{in(\omega_1 t + \phi_1)} \right] \quad (\text{A3})$$

$$= 2\pi h_0 e^{i\phi_0} \delta(\omega - \omega_0) \star \sum_{n=-\infty}^{\infty} e^{in\pi/2} e^{in\phi_1} J_n(\Gamma) F_t \left[ e^{in\omega_1 t} \right] \quad (\text{A4})$$

$$= 2\pi h_0 e^{i\phi_0} \delta(\omega - \omega_0) \star \sum_{n=-\infty}^{\infty} e^{in(\phi_1 + \pi/2)} J_n(\Gamma) \delta(\omega - n\omega_1) \quad (\text{A5})$$

$$= 2\pi h_0 e^{i\phi_0} \sum_{n=-\infty}^{\infty} e^{in(\phi_1 + \pi/2)} J_n(\Gamma) \delta(\omega - \omega_0) \star \delta(\omega - n\omega_1) \quad (\text{A6})$$

$$= 2\pi h_0 e^{i\phi_0} \sum_{n=-\infty}^{\infty} e^{in(\phi_1 + \pi/2)} J_n(\Gamma) \delta(\omega - \omega_0 - n\omega_1) \quad (\text{A7})$$

In this derivation we start by using the linearity of the Fourier Transform, then we invoke the Convolution Theorem, use the Jacobi-Anger Expansion, again use the linearity of the Fourier Transform, and the last few steps are simple Dirac Delta Function manipulations.

### Appendix B: The Fourier Transform of The Real-Valued Function

$$\mathcal{F}_t \left[ \Re \left\{ h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \right\} \right] = \frac{1}{2} \mathcal{F}_t \left[ h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} + h_0 e^{-i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \right] \quad (\text{B1})$$

$$= \frac{h_0}{2} \left[ e^{i\phi_0} \mathcal{F}_t \left[ e^{i(\omega_0 t + \Gamma \cos(\omega_1 t + \phi_1))} \right] + e^{-i\phi_0} \mathcal{F}_t \left[ e^{i(-\omega_0 t - \Gamma \cos(\omega_1 t + \phi_1))} \right] \right] \quad (\text{B2})$$

From the previous section we see that

$$\mathcal{F}_t \left[ e^{i(\omega_0 t + \Gamma \cos(\omega_1 t + \phi_1))} \right] = \sum_{n=-\infty}^{\infty} e^{in(\phi_1 + \pi/2)} J_n(\Gamma) \delta(\omega - \omega_0 - n\omega_1) \quad (\text{B3})$$

and by replacing  $\omega_0$  with  $-\omega_0$  and  $\Gamma$  with  $-\Gamma$ , and using the fact that  $J_n(-\Gamma) = (-1)^n J_n(\Gamma)$  we also have

$$\mathcal{F}_t \left[ e^{i(-\omega_0 t - \Gamma \cos(\omega_1 t + \phi_1))} \right] = \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \pi/2)} J_n(\Gamma) \delta(\omega + \omega_0 - n\omega_1) \quad (\text{B4})$$

so the Fourier transform of the real part of our frequency/phase modulated signal is

$$\begin{aligned} \mathcal{F}_t \left[ \Re \left\{ h_0 e^{i(\omega_0 t + \phi_0 + \Gamma \cos(\omega_1 t + \phi_1))} \right\} \right] &= \frac{1}{2} h_0 \left[ e^{i\phi_0} \sum_{n=-\infty}^{\infty} e^{in(\phi_1 + \pi/2)} J_n(\Gamma) \delta(\omega - \omega_0 - n\omega_1) \right. \\ &\quad \left. + e^{-i\phi_0} \sum_{n=-\infty}^{\infty} e^{in(\phi_1 - \pi/2)} J_n(\Gamma) \delta(\omega + \omega_0 - n\omega_1) \right] \\ &= \frac{1}{2} h_0 \sum_{n=-\infty}^{\infty} e^{in\phi_1} J_n(\Gamma) \left[ e^{i(\phi_0 + n\pi/2)} \delta(\omega - \omega_0 - n\omega_1) \right. \\ &\quad \left. + e^{-i(\phi_0 + n\pi/2)} \delta(\omega + \omega_0 - n\omega_1) \right] \end{aligned} \quad (\text{B5})$$

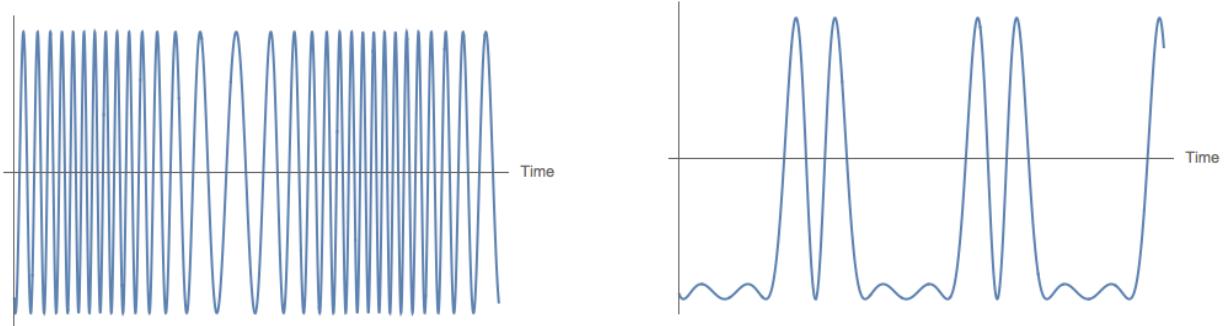
$$\hat{f}(\omega) = \frac{1}{2}h_0 \sum_{n=-\infty}^{\infty} e^{in\phi_1} J_n(\Gamma) \left[ e^{i(\phi_0+n\pi/2)} \delta(\omega - \omega_0 - n\omega_1) + e^{-i(\phi_0+n\pi/2)} \delta(\omega + \omega_0 - n\omega_1) \right] \quad (\text{B6})$$

### Appendix C: Intrinsic Mode Functions

Intrinsic mode functions are functions that are obtained by decomposing a signal. In general, intrinsic mode functions add up to the signal that was decomposed and are usually extracted by empirical mode decomposition [8].

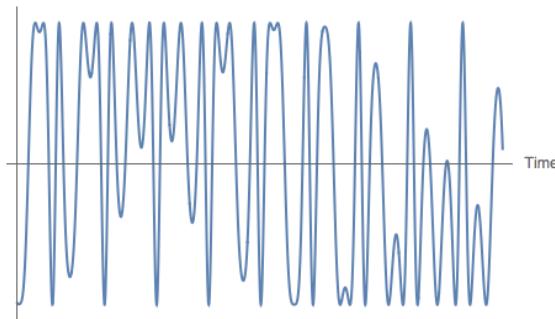
Classically, the extraction of intrinsic mode functions also yields a component named the residue — which is the back-bone of the signal. In this section, we shall assume the residue is one of the intrinsic mode functions for simplicity. Suppose we want to disintegrate a signal  $f(t)$  into  $k$  intrinsic mode functions. We first find the local maxima and minima of  $f(t)$ . Thereafter, we interpolate the local maxima to create an upper envelope and interpolate the local minima to create a lower envelope via cubic splines. The mean of these envelopes is computed and subtracted from the input signal to result in the first intrinsic mode function. This intrinsic mode function is then subtracted from the input signal to result in the residue. This approach is repeated  $k - 1$  times in total to obtain  $k$  intrinsic mode functions (where the last intrinsic mode function is the residue).

### Appendix D: Selected Waveforms



(a) When the carrier frequency is much higher than the modulation frequency the frequency modulated wave is "nice". One example of the exotic signals cyclic frequency modulation can produce.  $\Gamma = 9.88, \omega_0 = 3.62, \omega_1 = 0.21, \phi_1 = 3.14$

(b) When the carrier frequency is a multiple of the modulation frequency one can get exotic waveforms. One example of the exotic signals cyclic frequency modulation can produce.  $\Gamma = 2.88, \omega_0 = 1.5, \omega_1 = 0.75, \phi_1 = 0$



(c) When the carrier and modulation frequencies are almost equal the resulting FM wave is not so nice. This is a particularly nasty waveform that could be mistaken as noise by looking at it.  $\Gamma = 3, \omega_0 = 0.42, \omega_1 = 0.46, \phi_1 = 0.1$

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- [1] Junaid Aasi, J Abadie, BP Abbott, Richard Abbott, TD Abbott, MR Abernathy, Carl Adams, Thomas Adams, Paolo Addesso, RX Adhikari, et al. Enhanced sensitivity of the ligo gravitational wave detector by using squeezed states of light. *Nature Photonics*, 7(8):613–619, 2013.
- [2] A Abramovici, W Althouse, J Camp, D Durance, JA Giaime, A Gillespie, S Kawamura, A Kuhnert, T Lyons, FJ Raab, et al. Improved sensitivity in a gravitational wave interferometer and implications for ligo. *Physics Letters A*, 218(3):157–163, 1996.
- [3] Mina Aminghafari, Nathalie Cheze, and Jean-Michel Poggi. Multivariate denoising using wavelets and principal component analysis. *Computational Statistics & Data Analysis*, 50(9):2381–2398, 2006.
- [4] Albert Einstein. *On gravitational waves*. Springer, 1990.
- [5] Deanna Emery. A coherent three dimensional fft based search scheme for gravitational waves from binary neutron star systems. *LIGO DCC*, (T1500307-v2), September 2015.
- [6] B.P. Abbott et al. (LIGO Scientific Collaboration and Virgo Collaboration). Observation of gravitational waves from a binary black hole merger. *Physical Review Letters*, 116(061102):1–16, 12 February 2016.
- [7] Gerald B. Folland. *Fourier Analysis and its Applications*. Number ISBN 0-534-17097-3. Brooks and Cole Publishing Company, 1992.
- [8] Norden E Huang, Zheng Shen, Steven R Long, Manli C Wu, Hsing H Shih, Quanan Zheng, Nai-Chyuan Yen, Chi Chao Tung, and Henry H Liu. The empirical mode decomposition and the hilbert spectrum for nonlinear and non-stationary time series analysis. In *Proceedings of the Royal Society of London A: Mathematical, Physical and Engineering Sciences*, volume 454, pages 903–995. The Royal Society, 1998.
- [9] Nadja S Magalhães, Warren W Johnson, Carlos Frajula, and Odylio D Aguiar. Determination of astrophysical parameters from the spherical gravitational wave detector data. *Monthly Notices of the Royal Astronomical Society*, 274(3):670–678, 1995.
- [10] Karim G Oweiss and David J Anderson. Noise reduction in multichannel neural recordings using a new array wavelet denoising algorithm. *Neurocomputing*, 38:1687–1693, 2001.
- [11] Eric Poisson. Gravitational waves from inspiraling compact binaries: The quadrupole-moment term. *Physical Review D*, 57(8):5287, 1998.
- [12] Charles Rader. Discrete fourier transforms when the number of data samples is prime. *Proceedings of the IEEE*, 56(6):1107–1108, 1968.
- [13] Peter R. Saulson. *Fundamentals of Interferometric Gravitational Wave Detectors*. World Scientific Publishing Company, March 31, 1994.