

Revolutionary Quasar Physics with the JWST and MIRI MRS

Scientific Category: Massive Black Holes And Their Host Galaxies

Scientific Keywords: AGN Host Galaxies, High-Luminosity AGN/Quasars, M-Sigma Relation, Supermassive Black Holes, Winds and Outflows

Alternate Category: Galaxies and the IGM

Instruments: MIRI

Proprietary Period: 0 months

Allocation Information (in hours):

Science Time: 14.3

Charged Time: 22.2

Abstract

Several fundamental facts remain unknown for high-luminosity AGN, i.e. quasars: What is the main AGN triggering mechanism at the height of quasar activity at redshifts $z=2-3$? What is the direct, observational evidence in individual objects, that links AGN activity to star formation? Can we observe AGN 'feedback' in action and in situ for the most luminous sources at their peak activity? These remain among the outstanding observational extragalactic questions of our time. And they will be answered with the launch of the James Webb Space Telescope.

Here, we propose the perfect Director's Discretionary Early Release Science (DD ERS) program. We have identified a population of mid-infrared bright quasars at the peak of cosmological quasar activity, $z\sim 2.5$. Their global star-formation properties are currently unknown, but observations with JWST MIRI, in particular MRS spectroscopy, will quantify the level of star-formation in these objects. Moreover, the IFU aspect of the Medium Resolution Spectrometer will allow, for the first time, detailed investigations of the both the central AGN IR emission and any potentially extended emission.

We have already begun to design and create science-enabling products (SEPs) to help the community understand JWST's capabilities. Our MIRI MRS Code Repo, github.com/miri-mrs, is already active and completely accessible to anyone in the broader community. We will deliver the MIRI MRS SEPs first with mock data before the launch of JWST, and then in rapid fashion once the start of science operations commences in April 2019. This proposal, and all the accompanying code, calculations and documentation can be found at: github.com/d80b2t.

Investigators:

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	D Kocevski	Colby College	USA/ME
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*	M Ward	Durham Univ.	GBR
	N Zakamska	The Johns Hopkins University	USA/MD

Number of investigators: 22

* ESA investigators: 9

Target Summary:

Target	RA	Dec
J0834+0159	08 34 48.4800	+01 59 21.10
J1232+0912	12 32 41.7300	+09 12 9.30
J2215-0056	22 15 24.0000	-00 56 43.80
J2323-0100	23 23 26.1700	-01 00 33.10

Observing Summary:

Target	Observing Template	Flags	Allocation *
J0834+0159	MIRI Medium Resolution Spectroscopy		12,912 / 19,973
J1232+0912	MIRI Medium Resolution Spectroscopy		12,912 / 19,973
J2215-0056	MIRI Medium Resolution Spectroscopy		12,912 / 19,973
J2323-0100	MIRI Medium Resolution Spectroscopy		12,912 / 19,973

* Science duration / charged duration (sec)

Total Prime Science Time in Hours: 14.3

Total Charged Time in Hours: 22.2

Observing Description

We propose to use MIRI MRS to observe four IR-bright quasars across the full MRS wavelength range. We will deliver to the community Science Enabling Products (SEPs) that will enable the accumulation, reduction and timely analysis of MRS data, focusing on extra-galactic science.

Our 4 quasars are all well detected in the WISE W3 and W4 bands, and are thus very bright across 7.5-17um and 20-28um. They were first identified via their optical-to-IR color, initially via the r-W4 color, refined to an i-W3 color selection and are in the SDSS/BOSS quasar catalogs. As such, they have come to be known as "Extremely Red Quasars" (ERQs).

We will observe four of these Extremely Red Quasars (ERQs) with the MIRI Medium Resolution Spectrometer. All our targets are within 10 degrees of the Celestial Equator, and 3 out of our 4 targets are within 2 degrees. Our targets are at approximately 08h30m, 12h30m, 22h15m and 23h25min, in Right Ascension. We will observe each target for 19973 seconds in total, 12912 seconds of which are Science exposure. We will cover the full wavelength range, so will observe in the SHORT, MEDIUM and LONG modes of the MRS.

Taking advantage of the PI's host institution, and with continuing consultation with the MIRI team, as well as acknowledging the rationale of the ERS in general, we have decided to focus in on one instrument, and one particular mode of that instrument, and make sure we deliver to the (extragalactic) community the very best data tools, observational preparation instructions and science enabling products (SEPs). Our major documentation and notes are kept at::

github.com/d80b2t/JWST_ERS

with our Observing notes kept at::

https://github.com/d80b2t/JWST_ERS/tree/master/ETC_calcs

We encourage the Program Coordinators and Schedulers to investigate these links.

■ Rationale for DD-ERS Program

We have identified a population of obscured, mid-infrared bright quasars at the peak of cosmological quasar activity. These sources are mid-IR luminous and may be powered by major bursts of star formation tied to an early phase of galaxy evolution/formation. However, their global star-formation properties are currently unknown, but observations with JWST MIRI, and in particular MRS spectroscopy, will quantify the level of star-formation in these objects. *We will observe four “extremely red” quasars with MIRI MRS across the full wavelength range to very high signal to noise.* These observations will address the fundamental question of the link between star-formation and AGN activity, at the cosmological epoch for both of these processes; an investigation JWST was specifically built for.

Moreover, we have already begun to design and create science-enabling products (SEPs) to help the community understand JWST’s capabilities, Our **MIRI MRS Code Repo** is already active and completely accessible to anyone in the broader community. We will deliver the MIRI MRS SEPs first with mock data before the launch of JWST, and then in rapid fashion once the start of science operations commences in April 2019. Our team’s commitment to an open access ideology, not only for data, but for analysis codes, documentation, and scientific manuscripts is already evident and in place (for an example, see **the P.I.’s GitHub**). Via these code repos, and extensive documentation efforts, we aim to engage a broad cross-section of the astronomical community, in particular the extra-galactic community, in familiarizing themselves with JWST data and its scientific capabilities.


Science Rationale

Over 50 years after their formal identification, and over two decades since the calculation of their space density evolution, several fundamental facts remain unknown for high-luminosity AGN, i.e. quasars: What is the main AGN triggering mechanism at the height of quasar activity at redshifts $z = 2-3$? What direct, observational evidence in individual objects links AGN activity to star formation? Can we observe “AGN feedback” in action, in situ, for the most luminous sources at their peak activity? These remain the outstanding observational extragalactic questions of our time. And they will be answered with the launch of the *James Webb Space Telescope*.

The recently identified Extremely Red Quasar (ERQ) objects are a unique obscured quasar population with extreme physical conditions related to powerful outflows across the line-forming regions. These objects are found at the same cosmological epoch as the peak of quasar activity, $z \approx 2.5$, and are the best candidates to date to show outflows on galactic scales in high-luminosity objects at these epochs.

However, it is unknown whether the large IR luminosities observed in these quasars is from star formation, which should produce strong polycyclic aromatic hydrocarbon (PAH) spectral features, or if the hot dust near the central quasar, which should produce near-PAH, is driving the MIR luminosity. *Via the detection, or otherwise, of PAH spectral features, we will*


measure the SFRs during what is potentially a very early/obscured stage of massive galaxy formation in the Extremely Red Quasar population.


MIRI MRS is the instrument of choice since no other medium resolution spectrometer on JWST observes longward of $5\mu\text{m}$; going redder than this is crucial  order to detect PAH features in $z > 2$ objects. The desire to immediately gain high signal-to-noise spectra in order to investigate the physics and chemistry of quasar PAHs, along with observational overhead concerns, pushes us to observe each object for 3.6 hours, for a total program Charged Time of 22.20 hours.

Community Access Rationale

We will satisfy the Goals and Principles of the DD ERS program of ensuring open access to our datasets and analyses in support of the preparation of Cycle 2 proposals, and, engaging a broad cross-section of the astronomical community, in particular the extragalactic community, in familiarizing themselves with JWST data and scientific capabilities by utilizing modern analysis (code writing) repository tools and documentation that is standardized and publicly editable.

In particular, we aim to produce several key science products:

- `mrs_analyzer` A Python module for analyzing MRS data; 
- `mrsfringe` A Python module for mitigating MRS fringing issues;
- a suite of science results on quasar feedback and evolution.

Critically, we have already begun working closely with the MIRI team (due to the P.I.'s location at Edinburgh) and will continue to develop tools here for the MIRI MRS. We will produce of SEP analysis code and documentation. ~~Critically,~~ with a maximum of 6 months between the first ERS observations (April 2019) and the Cycle 2 GO Call for Proposals (September 2019), this will be too short for dissemination of our findings, novel techniques and science results in the traditional manner, i.e. via published journal articles. Moreover, ongoing updated versions of our analyses are envisaged to happen until right up to the Cycle 2 deadline. To solve these issues, we will fully employ the power of a code version repository system, in our case GitHub, to keep the community informed and updated with ~~or~~ SEPs. GitHub  has code versioning automatically built-in so proper referencing of e.g. technical notes is straight-forward.

■ Scientific Justification

One lasting scientific legacy of the *Hubble Space Telescope (HST)* will be the discovery of giant black holes at the centers of galaxies, confirming the longstanding theory of the “central engines” of quasars. One of the major surprises from the *Hubble* was the discovery

of a correlation between black hole mass and host galaxy properties.¹ This connection, causal or otherwise, may provide crucial clues to how and why these black holes formed and how their host galaxies evolved. *As of the launch of the James Webb Space Telescope (JWST), the question of how black holes affect their host galaxies is one of the outstanding questions in astrophysics.*

Observational and theoretical work now suggests that active galaxies and black holes are potentially linked to both the triggering, and “quenching”, of massive star formation. The link between massive galaxies and the central super-massive black holes (SMBHs) that seem ubiquitous in them is now thought to be vital to the understanding of galaxy formation and evolution, and huge observational and theoretical effort has been invested in trying to measure and understand the physics involved in these systems. The two main energy sources available to a galaxy are nuclear fusion in stars and gravitational accretion onto compact objects, and we still do not fully understand the interaction of ~~an~~ active galactic nuclei and the star formation properties of their host galaxies, *especially at the epoch of maximal cosmic SFR and quasar activity, redshifts $z \sim 1 - 3$.*



The “quenching” of galaxy-wide star formation, is supposedly driven by “AGN feedback”, where the AGN helps to heat the surrounding gas corona, offsetting cooling losses and disrupting the gas inflow. This feedback manifests itself as high-velocity outflows from the AGN/quasar. However, *direct observational evidence* for AGN feedback is conspicuous by its absence. This is especially true at high- z , e.g. $z = 2 - 3$, at the height of the Quasar Epoch. We have identified the best candidates that suggest we are seeing quasar feedback in action, in situ at high-redshift. These are the “Extremely Red Quasars” identified via their WISE W3/4 colors. As such, these mJy luminous AGN *are ideal targets for JWST MIRI.*

The Extremely Red Quasar Population

ERQs are a unique obscured quasar population with extreme physical conditions related to powerful outflows across the line-forming regions. These sources are the signposts of the most extreme form of quasar feedback at the peak epoch of galaxy formation, and may represent an active “blow-out” phase of quasar evolution.

By matching the quasar catalogues of the Sloan Digital Sky Survey (SDSS), the Baryon Oscillation Spectroscopic Survey (BOSS) to the Wide-Field Infrared Survey Explorer (WISE), Ross et al. (2015) discovered quasars with extremely red infrared-to-optical colors: $r_{\text{AB}} - W4_{\text{Vega}} > 14$ mag, i.e., $F_{\nu}(22\mu\text{m})/F_{\nu}(r) \gtrsim 1000$; see Figure 1. These objects have infrared luminosities $\sim 10^{47}$ erg s⁻¹. The original motivation ~~here~~ was to look for PAHs that had redshifted in the WISE W4 band for $z \approx 2.5$ quasars in order to study and link luminous AGN activity and star formation at the “height of the quasar epoch”.

¹Assessment of Options for Extending the Life of the Hubble Space Telescope: Final Report (2005); <https://www.nap.edu/read/11169/chapter/5>).

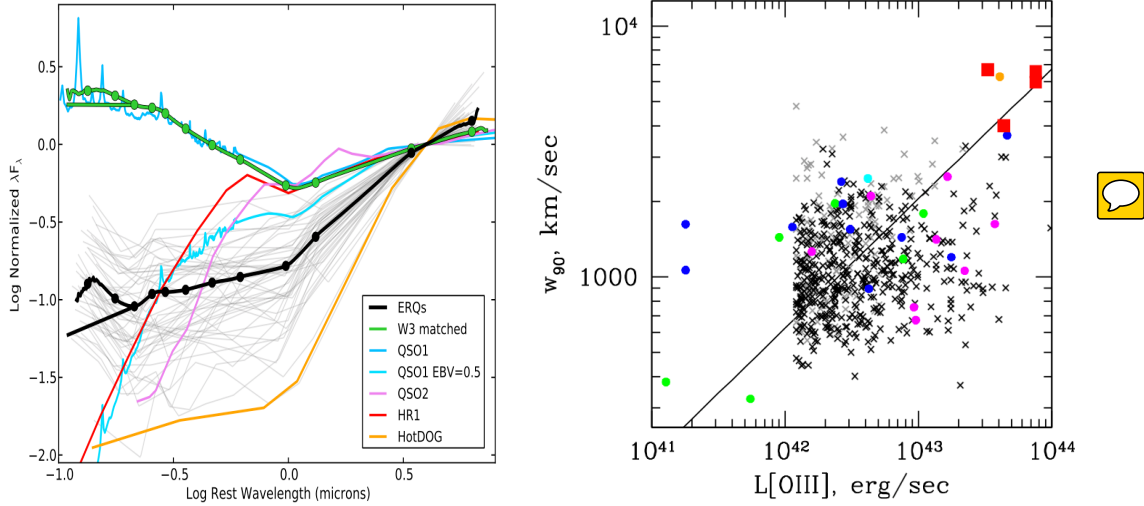


Figure 2: (Left) From Hamann et al., 2017, the normalized median SEDs for Type 1 non-BALs in the core ERQ sample (black curve) plus blue quasars matched to the core ERQs in W3 magnitude (green curve). The Type 1 quasar template with and without reddening equal to $E(BV) = 0.5$ (light blue, ‘QSO1’, from Polletta et al. 2007), a typical Type 2 quasar from Mateos et al. (2013; ‘QSO2’, purple), a typical heavily reddened Type 1 quasar from Banerji et al. (2013; ‘HR1’), and a typical HotDOG from Tsai et al. (2015). The light grey curves are SEDs of individual core ERQs. (Right) O III kinematics as a function of luminosities for our four targets, red squares, with other quasar populations are shown by black points and various colored symbols. At these extreme velocities, the gas cannot be confined by any realistic galaxy potential and is thus likely to escape from the galaxy; these objects are likely signposts of the extreme blow-out phase of quasar feedback observed at the peak epoch of galaxy formation.

of four $z \sim 2.5$ extremely red quasars (Zakamska et al. 2016). We discovered very broad (full width at half max = $2600 - 5000 \text{ km s}^{-1}$), strongly blue-shifted (by up to 1500 km s^{-1}) O III $\lambda 5007\text{\AA}$ emission lines in these objects. In a large sample of type 2 and red quasars, O III kinematics are positively correlated with infrared luminosity, and the four objects in our sample are on the extreme end both in O III kinematics and infrared luminosity. As such, we estimate that at least 3% of the bolometric luminosity in these objects is being converted into the kinetic power of the observed wind. Photo-ionization estimates suggest that the O III emission might be extended on a few kpc scales, which would suggest that the extreme outflow is affecting the entire host galaxy of the quasar.

Given the extreme nature of these objects, especially in the kinematics, what is actually powering the IR luminosity? Is it young, hot stars in a starburst, or hot dust powered by the central AGN? Furthermore, how, if at all, is this IR emission connected to the conditions at large of the quasar and host galaxy?

MIR spectroscopy: In order to quantify the main energy contribution in galaxies, one needs to differentiate objects powered by nuclear fusion in stars (i.e. normal and starburst galaxies) from objects powered by mass accretion onto supermassive black holes (quasars and AGNs). MIR spectroscopy allows the detection of obscured AGN components even when the MIR is dominated by the host galaxy. Several classification diagrams have been developed to determine the AGN contribution to an observed spectrum based on certain

spectral features, such as high ionisation emission lines like Ne v , [Ne II] and O IV, the EW of PAH features and the strength of the silicate feature at $9.7 \mu\text{m}$ (see, e.g. Spoon et al., 2007; Armus et al., 2007; Veilleux et al., 2009; Hernan-Caballero & Hatziminaoglou, 2011). A number of techniques have also been developed to model the observed MIR spectra and constrain the AGN and starburst contributions (see e.g. Schweitzer et al., 2008; Nardini et al., 2008; Deo et al., 2009; Feltre et al., 2013).



Polycyclic Aromatic Hydrocarbons: Polycyclic Aromatic Hydrocarbons (PAHs) are abundant, ubiquitous, and a dominant force in the interstellar medium of galaxies (see e.g., Tielens, 2008, ARAA, 46, 289 for a review). Aromatic features are already a significant component of dusty galaxy spectra as early as $z \approx 2$ (Yan et al., 2005, ApJ, 628, 604). and the infrared (IR) emission features at 3.3, 6.2, 7.7, 8.6, and $11.3 \mu\text{m}$ are generally attributed to IR fluorescence from (mainly) far-ultraviolet (FUV) pumped large ~~polycyclic aromatic hydrocarbon~~ (PAH) molecules. As such, these features trace the FUV stellar flux and are thus a measure of star formation (Peeters et al, 2004, ApJ, 613, 986). The interstellar IR emission spectrum is incredibly rich and shows a wealth of detail. It is dominated by major emission features at 3.3, 6.2, 7.7, 8.6, 11.2, 12.7, and $16.4 \mu\text{m}$. In addition, there are weaker features at 3.4, 3.5, 5.25, 5.75, 6.0, 6.9, 7.5, 10.5, 11.0, 13.5, 14.2, 17.4, and $18.9 \mu\text{m}$. Given the redshift of our ERQs and the MIRI wavelength coverage we will coverage $1.36 \leq \lambda_{\text{emitted}} \leq 8.6 \mu\text{m}$.

A final science goal, is the aim of obtaining *spatial information* you might get for the lines and PAHs. The IFU aspect of the Medium Resolution Spectrometer will allow, for the first time, detailed investigations of the both the central AGN IR emission and any potentially extended emission.

1 Bibliography

Fabian, 2012, ARAA 50, 455 • Kormendy & Ho, 2013, ARAA, 51, 511 • Yuan & Narayan, 2014, ARAA, 52, 529 • Heckman & Best, 2014, ARAA, 52, 589 •

■ Technical Description

We propose to use MIRI MRS to observe four quasars across the full MRS wavelength range in order to deliver to the community Science Enabling Products (SEPs) related to the accumulation, reduction and timely analysis of the MIRI Medium Resolution Spectrometer. Details of our four quasar targets, including the current state of our extensive multiwavelength follow-up programs, are given in Table 1.

Overall Experimental Design: We have honed in on a specific early release science case that will engage a broad cross-section of the astronomical community, and deliver a dataset for MIRI MRS that can be used for e.g. extragalactic galaxy and AGN studies in Cycle 2. After discussion with the MIRI Team (A. Glasse; priv. comm.), we settled on the ideology of picking out one instrument (MIRI) and one observing mode (MRS) and making sure we deliver the highest quality data analysis and SEPs here for the community. Observations with MIRI MRS will also directly answer the science questions we have posed.

MIRI MRS is the instrument of choice since no other medium resolution spectrometer on JWST observes longward of 5 μ m; going redder than this is crucial in order to detect PAH features in z \sim 2 objects. The desire to immediately gain high signal-to-noise spectra in order to investigate the physics and chemistry of quasar PAHs, along with observational overhead concerns, pushes us to observe each object for 3.6 hours, for a total program Charged Time of 22.20 hours.

MIRI MRS Observing Overview

We will be utilizing the medium-resolution integral field unit (IFU) and spectroscopy mode. For our MRS operations we shall operate over the full wavelength range; employ the 4-point dithering pattern, optimized for point sources and using the SLOW detector read-out mode. For the observations themselves, a wide range of considerations (sample size, desired very high SNR, time available for ERS programs, level of read noise etc.) we converge on: 15 Groups 3 Integrations and 1 Exposures for a total of 12,912 seconds science exposure per object. *Using Smart Accounting, the total Charged Time is 22.20 hours.*



MIRI Observing Details

Our observational set-up is:

- MIRI MRS.
- Full spectral coverage; thus we will use all three different spectral settings, SHORT (A), MEDIUM (B), and LONG (C).
- Since with the ERQs we are likely to be observing either point sources or compact sources, we choose to use the point source optimized, “4-point ALL” dither pattern

Object Name (SDSS)	J0834+0159	J1232+0912	J2215-0056	J2323-0100
Object R.A.	08:34:48.48	12:32:41.73	22:15:24.00	23:23:26.17
object declination	+01:59:21.1	+09:12:09.3	−00:56:43.8	−01:00:33.1
r -band AB magnitude	21.20±0.05	21.11± 0.05	22.27±0.12	21.62± 0.08
WISE W4-band Vega magnitude	6.88±0.09	6.78 ±0.09	7.91±0.24	7.76±0.22
WISE W4-band flux, F_ν	>6 mJy	>6 mJy	6 mJy	>6 mJy
$i_{\text{AB}} - W3_{\text{AB}}$	6.0	6.8	6.2	7.2
Redshift z	2.591	2.381	2.509	2.356
REW C IV	209±6	225±3	153±5	256±5
C IV FWHM km s ^{−1}	2863±65	4787±52	4280±112	3989±62
O III FWHM erg s ^{−1}	2811	4971	3057	2625
Spectro-polarimetry	×	✓	✓	×
VLA data	?	?	?	?
ALMA Band 6	tbc	✓	tbc	✓
<i>HST</i> Cycle 24	ACS/WFC3 obtained	ACS/WFC3 <i>pending</i>	ACS/WFC3 <i>pending</i>	ACS/WFC3 <i>pending</i>
JWST target visibility (Start)	2019-04-01	2019-05-08	2019-05-22	2019-06-07
JWST target visibility (End)	2019-05-07	2019-07-01	2019-07-15	2019-07-29

Table 1: Our four Extremely Red Quasar targets. All four quasars were first identified in Ross et al. (2015). Values of e.g. REWs, FWHMs are from Zakamska et al. (2016) and Hamann et al. (2017).

(4PTALL dither). According to the pre-flight expected relative performance of MRS dither patterns (Figure 3, <https://jwst-docs.stsci.edu/display/JTI/MIRI+MRS+Dithering>) this is the only MIRI MRS dither patterns that guarantees “GOOD” (i.e. half-integer) sampling throughout the common field of view across all four channels. We opt for the 4-point ALL dither pattern, point source optimized. As noted in the preparation pages, this provides robust performance at all wavelengths and adequate point source separation in all channels such that dedicated background observations are not required. It is also the only dither where spatial sampling is “GOOD” (i.e. half-integer sampling) throughout the common field of view across all four channels.

- We are interested in only a single object per point, so no mosaicking is necessary.
- MIRI Detector Readout mode: SLOW JWST MIRI’s “Slow mode” readout pattern offers fewer detector artifacts and slightly lower detector noise than the “fast mode”, making it a good choice for faint source medium-resolution spectroscopy where the sky backgrounds are very low. This is exactly what we want for our ERQ observations.
- Subarray is FULL. This is fixed for MRS.

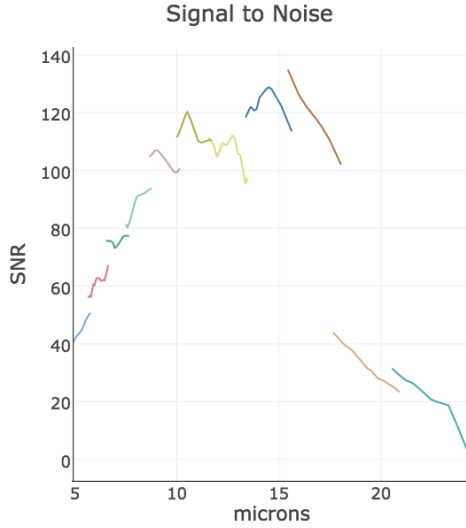


Figure 3: The evolution of linear bias, b , for SDSS Quasars (Ross et al. 2009). We find that quasars tend to inhabit haloes of constant mass, $M_{\text{DM}} \sim 2 \times 10^{12} M_{\odot}$ from $z \sim 3$ to the present day.

We take the “core ERQ” SED that is given in Hamann et al. (2017) and is fully representative of the ERQ population at large. The file ~~(found on the GitHub)~~ of `core_ERQ_SED_notLog.dat` is used here. We normalize this SED at a wavelength of $23\mu\text{m}$ to a source flux density of 5mJy , again very representative of the WISE W3/4-detected ERQ population given in Ross et al and Hamann et al.

At this stage we *do not* include any emission (or absorption) lines since... In the ETC we assume the Shape of the source is Point. Other notes, using the ETC, include having: Medium Backgrounds; ‘IFU Nod In Scene’; Aperture location Centered on source; Aperture radius of $0.3''$; Nod position in scene of $X = Y = 0.5''$. Our JWST ETC Workbook has `wb` ID 7474.

MIRI MRS Target Acquisition

Observations with the MIRI MRS IFU may often require a target acquisition (TA) procedure, especially at the shortest wavelengths.

The required pointing precision for the MIRI MRS is 90 mas (1σ radial), which is approximately the half width of a slice at the shortest wavelength. This required accuracy limits the spacecraft move between the TA region in the imager and the center of the IFU to $\pm 50''$ (Sivaramakrishnan et al. 2006). The TA may be achieved with the FND, F560W, F1000W and F1500W on either the target or a suitable offset that is $\pm 50''$. The TA centroiding procedure loses accuracy if the pixels are saturated so a brightness limit (Table 1) must be considered for the target.

APT, Overheads and Smart Accounting.

All the above details in are out APT. Visit Planner was loaded and Smart Accounting run. Our full proposal comes in at 20.32 hours of charged time.

■ **Plan for Alternative Targets**

We have 15 ERQs that will be observed with ALMA Cycle 5, Band 6, four of which are our primary targets. The remaining 11 objects are given here as alternative targets. As can be seen, redshifts and e.g. WISE W3/W4-band magnitudes and fluxes are essentially the same for the core targets, hence we do not consider changing the observational strategy.

■ **Special Requirements**

There are no Special Observational Requirements for our ERS.

■ **Justify Coordinated Parallel Observations**

We are not inducing Coordinated Parallels observations.

■ **Justify Duplications**

There are no duplicated observations.

■ Data Processing & Analysis Plan

General Outline and Motivations Driving Data Processing & Analysis

We are already ensuring open access to representative data sets, processing pipelines and analysis tools in support of the preparation of both Cycle 1 and Cycle 2 proposals. The key links are::

`github.com/d80b2t/JWST_ERS`
(the P.I.'s personal and public GitHub repository).

`github.com/miri-mrs`
(a “Organizational” public GitHub repository for the Community).

`miri-mrs.readthedocs.io`
(With the source code here: <https://github.com/miri-mrs/ERQs/tree/master/docs>).

Our ERQ ERS proposal is the first part of a multi-cycle proposal project and plan. As such, we are *already highly motivated to produce the data processing tools, codes, documentation and identify the critical science-enabling products well in advance of the release of the Cycle 2 Call for Proposals (September 2019)*.

Our Science-Enabling Products team consists of five main parts::

- “Core Coders”;
- Website and ReadTheDocs Writers;
- Observational Follow-up;
- Senior members of staff, able to supply students/ ask for funding;
- A small steering committee.

ERS ERQ Science Enabling Products

`mrs_analyzer` is a Python module for analyzing MRS data;

`mrsfringe` is a Python module for mitigating MRS fringing issues;

And will integrate with...

<https://github.com/STScI-JWST>

see e.g.

https://github.com/STScI-JWST/jwst/blob/master/jwst/mrs_imatch/mrs_imatch_step.py

MIRISim

MIRI data simulations (at ESA) include an Integral Field observation with the Medium Resolution Spectrograph (MRS), a Low Resolution Spectrograph (LRS) observation, and an imaging observation. (Credit: ESA, Pamela Klaassen and the MIRISim Team).

Delivery Schedule for Science-Enabling Products.

Proposals must present a delivery schedule for science-enabling products. A description of STScI pipeline data products, processing and analysis software, and their anticipated availability, will be provided by the May 2017 release of the final version of this Call for Proposals. Proposers may consider multiple deliveries, with more advanced products provided over longer timescales. Proposals may include the collection, processing and analysis of ancillary data as part of an integrated DD ERS proposal.

Co-Investigators and Delivery of Science-Enabling Products.

Co-Investigators, together with the PI (and any Co-PIs) comprise a core team with the responsibility of developing and delivering science-enabling products as described in the proposal, as well as carrying out selected key aspects of the science investigations. A Co-I must have a well-defined, and generally sustained, continuing role in team activities, serve under the direction of the PI, or co-PI(s). Co-investigators may or may not receive funding, pending eligibility, through the DD ERS program. A bibliography and SEP tasks of our team can be found [here](#) and the SEP Deliverables are [here](#).