

Near and Mid-infrared properties of known $z \geq 5$ Quasars

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ABSTRACT

In this paper, we, for the first time since the discovery of $z \geq 5$ quasars, assemble all spectroscopically confirmed very high redshift quasars in one catalogue. In particular we present the near ($zZyYJHK_s$ and K) infrared and mid-infrared (WISE) properties of all 424 spectroscopically confirmed redshift $z \geq 5.00$ quasars. Using archival public WFCAM/UKIRT and VIRCAM/VISTA data we check for photometric variability in the near-infrared that might be expected from Super-Eddington accretion and find *blah*. We present a comprehensive series of colour-redshift and colour-colour plots and make inferences into the hot dust properties of the very high-redshift quasar population. Extrapolating the known quasar luminosity function we suggest that $x\%$ of the possibly detected $z \geq 5$ quasars in the current datasets have been discovered.

Key words: Astronomical data bases: surveys – Quasars: general – galaxies: evolution – galaxies: infrared.

1 INTRODUCTION

Very high redshift quasars (VHzQ; defined here to have redshifts $z \geq 5.00$) are excellent probes of the early Universe. This includes studies of the Epoch of Reionization for hydrogen (see e.g. Fan et al. 2006; Mortlock 2016, for reviews), the formation and build-up of supermassive black holes (e.g., Rees 1984; Wyithe & Loeb 2003; Volonteri 2010; Agarwal et al. 2016; Valiante et al. 2018; Latif et al. 2018) and early metal enrichment (see e.g., Simcoe et al. 2012; Chen et al. 2017; Bosman et al. 2017).

Super-Eddington (also called super-critical) accretion is a viable physical mechanism to explain the high luminosity and rapid growth of supermassive black holes in the early universe (e.g., Alexander & Natarajan 2014; Madau et al. 2014; Volonteri et al. 2015; Pezzulli et al. 2016; Lupi et al. 2016; Pezzulli et al. 2017; Takeo et al. 2018). Thus, one might well expect VHzQs to vary in luminosity as they potentially go through phases of Super-Eddington accretion and these signatures of photometric variability should be looked for, especially when the e.g. near-infrared K -band is sampling rest-frame $\approx 3670, 3145, 2750\text{\AA}$ at redshifts $z = 5, 6$ and 7 , respectively.

Quasars are also known to be prodigious emitters of infrared emission, thought to be from the thermal emission of dust grains heated by continuum emission from the accretion disc (e.g., Richards et al. 2006; Leipski et al. 2014; Hill et al. 2014; Hickox et al. 2017). Observations in the mid-infrared,

e.g. $\sim 3\text{--}30\mu\text{m}$ allow discrimination between AGN¹ and passive galaxies due to the $1.6\mu\text{m}$ “bump” entering the MIR at $z \approx 0.8 - 0.9$ (e.g., Wright et al. 1994; Sawicki 2002; Lacy et al. 2004; Stern et al. 2005; Richards et al. 2006; Timlin et al. 2016) as well as AGN and star-forming galaxies due to the presence of Polycyclic Aromatic Hydrocarbon (PAHs) at $\lambda > 3\mu\text{m}$ (e.g., Yan et al. 2007; Tielens 2008).

Jiang et al. (2006) and Jiang et al. (2010) report on the discovery of a quasar without hot-dust emission in a sample of 21 $z \approx 6$ quasars. Such apparently hot-dust-free quasars have no counterparts at low redshift. Moreover, those authors demonstrate that the hot-dust abundance in the 21 quasars builds up in tandem with the growth of the central black hole. Thus understanding how dust first forms and appears in the central engine remains an open question.

WISE mapped the sky in 4 passbands, in bands centered at wavelengths of $3.4, 4.6, 12,$ and $23\mu\text{m}$. In total the release all sky “ALLWISE” catalog, contains nearly 750 million detections at high-significance². Assef et al. (2013), Stern et al. (2012). Blain et al. (2013) presented WISE mid-infrared (MIR) detections of 17 (55%) of the then known 31 quasars at $z > 6$. However, Blain et al. (2013) was compiled with the WISE ‘All-Sky’ data release, as opposed to the superior “AllWISE” catalogs. That sample only examined the

¹ Historically, “quasars” and “Active Galactic Nuclei (AGN)” have described different luminosity/classes of objects, but here we use these terms interchangeably (with a preference for quasar) in recognition of the fact that they both describe accreting supermassive black holes (e.g. Haardt et al. 2016).

² wise2.ipac.caltech.edu/docs/release/allwise/expsup/sec2.1.html

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31 known $z > 6$ quasars; our sample has 148 objects with redshift $z \geq 6.00$.

Here we update Jiang et al. (2010) and Blain et al. (2013) (along with Table 8 of Bañados et al. 2016). Our motivations are numerous and include: (i) establishing the first complete catalogue of $z > 5.00$ quasars since the pioneering work from SDSS; (ii) analysing all the WFCAM and VISTA [NJC: homogeneous isn't the write term. The NIR data is very heterogeneous] near-infrared photometry for the quasars; (iii) making the first study of NIR variability of the VHzQ population and (iv) establishing the photometric properties for upcoming surveys and telescopes, e.g. the Large Synoptic Survey Telescope (LSST)³, ESA *Euclid*⁴ and the *James Webb Space Telescope* (JWST)⁵

This paper can be considered an update of Blain et al. (2013) and also an extension of parts of Bañados et al. (2016), with the latter study reporting WISE W1, W2, W3 and W4 magnitudes for the Panoramic Survey Telescope and Rapid Response System 1 (Pan-STARRS1, PS1; Kaiser et al. 2002, 2010), but with no further investigation into the reddest WISE waveband for the VHzQs. Bañados et al. (2016) reports and investigates the W1, W2 and W3 properties of quasars at $z > 5.6$. We chose redshift $z = 5.00$ as our lower redshift limit due to a combination of garnishing a large sample, adequately spanning physical properties (e.g. luminosity, age of the Universe) and to incorporate what knowledge we have gained over the last couple of decades since $z > 5$ quasars were discovered.

This paper is organized as follows. In Section 2, we present the assembled list of the 424 $z \geq 5.00$ VHzQs that we have compiled. We then give a high-level overview of the photometric surveys and datasets we use and present the photometry of the VHzQs. In Section 3 we In Section 4 we and we then note our Conclusions in Section 5.

We make the somewhat unconventional decision [NJC: Is this unconventional?] to present all our photometry and magnitudes on the AB zero-point system (Oke & Gunn 1983; Fukugita et al. 1996). This includes the near-infrared, as well as the mid-infrared magnitudes. Because of established conventions, we report SDSS [NJC: PanSTARRS/DES rather than SDSS?] *ugriz* magnitudes on the AB zero-point system (Oke & Gunn 1983; Fukugita et al. 1996), while the WISE W1–4 magnitudes are calibrated on the Vega system (Wright et al. 2010). For WISE bands, $m_{AB} = m_{Vega} + m$ where $m = (2.699, 3.339, 5.174, 6.66)$ for W1, W2, W3 and W4, respectively (Cutri et al. 2011; Brown et al. 2014). WFCAM and VISTA (?) are also calibrated on a Vega system (WFCAM), or near Vega system (VISTA). Brown et al. (2014), in PASA, is the paper about Recalibrating the Wide-field Infrared Survey Explorer (WISE) W4 Filter. We make use of the Explanatory Supplement to the WISE All-Sky Data Release, as well as the WISE AllWISE Data Release Products online. We use a flat Λ CDM cosmology with $H_0 = 67.7$ km s^{−1} Mpc^{−1}, $\Omega_M = 0.307$, and $\Omega_\Lambda = 0.693$ (Planck Collaboration et al. 2016) in order to be consistent with Bañados et al. (2016).

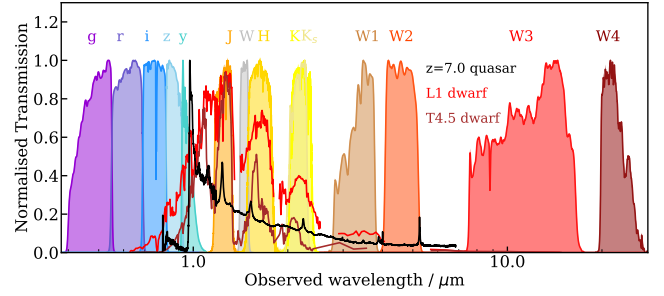


Figure 1. The spectral bands used by different survey telescopes and that are relevant here. The *grizy* filters are from the Pan-STARRS survey. The *JHK* are from UKIRT/WFCAM, while *K_s* is a VISTA/VIRCAM filter. The narrow W-band centered at $\lambda \approx 14,500\text{\AA}$ is a CFHT/Wircam filter. [NJC: Do we use the W-band anywhere?] The WISE passbands W1–4 are also presented. The quasar spectrum is a composite based on Vanden Berk et al. (2001) and Bañados et al. (2016). The L and T dwarf spectra are from Cushing et al. (2006).

2 DATA

In Table 1 we present our dataset and we have assembled this in the following manner. First we compile the list of all known, spectroscopically confirmed quasars from the literature. Most of these objects are easily identified by their broad Ly α emission line, N V emission and characteristic shape blueward of 1215\AA observed. As we shall see, some of the more recently discovered objects are close to the galaxy luminosity function characteristic luminosity M^* , and some have relatively weak or maybe even completely absorbed Ly α (e.g. Figures 7 and 10 in Bañados et al. 2016). We leave aside detailed investigation and discussion into spectral features and line strengths, and take as given the published spectra.

We then obtain optical, near-infrared and mid-infrared photometry for the spectral dataset. The optical data comes from the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS) survey (Chambers et al. 2016) and the Dark Energy Camera (Flaugher et al. 2015), including the Dark Energy Survey (DES; Flaugher 2005) and the Dark Energy Camera Legacy Survey (Dey et al. 2018, DECaLS).

Recently, Hyper Suprime-Cam (HSC; Miyazaki et al. 2018) has embarked on the Subaru Strategic Program (SSP; Aihara et al. 2018b) delivering exquisite multi-band photometry data (Aihara et al. 2018a) the five broad bands *grizy*.

The near-infrared data comes from two primary sources. First, the Wide Field Camera (WFCAM; Casali et al. 2007) on the the United Kingdom Infra-Red Telescope (UKIRT) primarily, but not exclusively as part of the UKIRT Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007). And second, data from the VIRCAM (VISTA InfraRed CAMera) on the VISTA (Visible and Infrared Survey Telescope for Astronomy; Emerson et al. 2006; Dalton et al. 2006).

We also obtain mid-infrared, $\lambda = 3 - 30\mu\text{m}$ wavelength data from the the Wide-Field Infrared Survey Explorer (WISE; Wright et al. 2010; Cutri 2013) mission. WISE mapped the sky in 4 passbands, in bands centered at wavelengths of 3.4, 4.6, 12, and $23\mu\text{m}$. In total the release all sky

³ lsst.org

⁴ sci.esa.int/euclid/

⁵ jwst.nasa.gov; sci.esa.int/jwst; www.asc-csa.gc.ca/eng/satellites/jwst; jwst.stsci.edu

“ALLWISE” catalog, contains nearly 750 million detections at high-significance⁶.

Figure 1 displays the wavelength and normalised transmission of the filters in question.

⁶ wise2.ipac.caltech.edu/docs/release/allwise/

| na | desig | ra_hms | dec_dms | ra | dec | redshift | mag | M1450 | ref |
|--------|-------------------|-------------|--------------|----------|-----------|----------|-------|--------|----------------|
| PSO | J000.3401+26.8358 | 00:01:21.63 | +26:50:09.17 | 0.340113 | +26.83588 | 5.75 | 19.52 | -27.16 | 1/1/1 |
| SDSS | J0002+2550 | 00:02:39.39 | +25:50:34.80 | 0.664117 | +25.84304 | 5.82 | 19.39 | -27.31 | 5/22/1 |
| SDSS | J0005-0006 | 00:05:52.34 | -00:06:55.80 | 1.468083 | -00.11549 | 5.85 | 20.98 | -25.73 | 5/12/1 |
| PSO | J002.1073-06.4345 | 00:08:25.77 | -06:26:04.60 | 2.107390 | -06.43456 | 5.93 | 20.41 | -26.32 | 1;43/1/1 |
| SDWISE | J0008+3616 | 00:08:51.43 | +36:16:13.49 | 2.214292 | +36.27041 | 5.17 | 19.12 | -27.34 | Wang2016 |
| PSO | J002.3786+32.8702 | 00:09:30.89 | +32:52:12.94 | 2.378702 | +32.87026 | 6.1 | 21.13 | -25.65 | 1/1/1 |
| SDSS | J0017-1000 | 00:17:14.68 | -10:00:55.4 | 4.311166 | -10.01540 | 5.011 | 99.99 | -99.99 | DR7.W16 |
| PSO | J004.3936+17.0862 | 00:17:34.47 | +17:05:10.70 | 4.393614 | +17.08631 | 5.8 | 20.69 | -26.01 | 1/1/1 |
| PSO | J004.8140-24.2991 | 00:19:15.38 | -24:17:56.98 | 4.814080 | -24.29920 | 5.68 | 19.43 | -27.24 | 1/1/1 |
| VDES | J0020-3653 | 00:20:31.46 | -36:53:41.8 | 5.131124 | -36.89495 | 6.9 | 99.99 | -99.99 | DES-VHS_inprep |

Table 1. All 424 $z \geq 5.00$ quasars that have been spectroscopically confirmed as of 2018 June. The first ten objects are given here as guidance to the format of the data table. The full table can be found online.

| na | desig | ra | dec | w1mag | w1err | w1snr | w2mag | w2err | w2snr | w3mag | w3err | w3snr | w4mag | w4err | w4snr |
|--------|-------------------|------------|--------------|--------|-------|-------|--------|-------|-------|--------|-------|-------|-------|-------|-------|
| PSO | J000.3401+26.8358 | 0.34011348 | 26.83588138 | 16.373 | 0.066 | 16.5 | 15.266 | 0.107 | 10.2 | 12.594 | 0.492 | 2.2 | 8.756 | -9.99 | 1.1 |
| SDSS | J0002+2550 | 0.66411726 | 25.84304425 | 16.162 | 0.057 | 19.0 | 15.542 | 0.127 | 8.5 | 12.416 | 0.423 | 2.6 | 8.683 | -9.99 | 1.2 |
| SDSS | J0005-0006 | 1.4680833 | -0.1154999 | 17.299 | 0.16 | 6.8 | 17.043 | -9.99 | 0.2 | 12.445 | -9.99 | -1.1 | 9.008 | -9.99 | -0.3 |
| PSO | J002.1073-06.4345 | 2.10739 | -6.43456 | 16.809 | 0.107 | 10.1 | 15.684 | 0.141 | 7.7 | 11.892 | -9.99 | 1.5 | 8.759 | -9.99 | 0.2 |
| SDWISE | J0008+3616 | 2.2142917 | 36.2704138 | 16.045 | 0.052 | 20.7 | 15.373 | 0.092 | 11.8 | 12.043 | -9.99 | 1.8 | 8.786 | -9.99 | 1.1 |
| PSO | J002.3786+32.8702 | 2.37870183 | 32.87026179 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -9.99 | -9.99 | -9.9 |
| SDSS | J0017-1000 | 4.3111666 | -10.01539722 | 15.936 | 0.055 | 19.7 | 15.167 | 0.094 | 11.5 | 12.026 | 0.334 | 3.2 | 8.52 | -9.99 | 1.2 |
| PSO | J004.3936+17.0862 | 4.39361347 | 17.08630447 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -9.99 | -9.99 | -9.9 |
| PSO | J004.8140-24.2991 | 4.81408 | -24.29916 | 16.281 | 0.069 | 15.8 | 15.569 | 0.116 | 9.4 | 12.123 | 0.344 | 3.2 | 8.82 | -9.99 | 0.5 |
| VDES | J0020-3653 | 5.1311237 | -36.8949476 | 16.844 | 0.094 | 11.6 | 16.354 | 0.204 | 5.3 | 12.679 | -9.99 | -0.1 | 8.342 | -9.99 | 0.8 |

Table 2. The mid-infrared photometric properties from the WISE ALLWISE catalogue for the 424 very-high redshift quasars. The first ten objects are given here as guidance to the format of the data table. The full table can be found online. *This is the third table here; the SECOND table is this with the NIR data...*

| na | desig | ra | dec | w1mag | w1err | w1snr | w2mag | w2err | w2snr | w3mag | w3err | w3snr | w4mag | w4err | w4snr |
|--------|-------------------|------------|--------------|--------|-------|-------|--------|-------|-------|--------|-------|-------|-------|-------|-------|
| PSO | J000.3401+26.8358 | 0.34011348 | 26.83588138 | 16.373 | 0.066 | 16.5 | 15.266 | 0.107 | 10.2 | 12.594 | 0.492 | 2.2 | 8.756 | -9.99 | 1.1 |
| SDSS | J0002+2550 | 0.66411726 | 25.84304425 | 16.162 | 0.057 | 19.0 | 15.542 | 0.127 | 8.5 | 12.416 | 0.423 | 2.6 | 8.683 | -9.99 | 1.2 |
| SDSS | J0005-0006 | 1.4680833 | -0.1154999 | 17.299 | 0.16 | 6.8 | 17.043 | -9.99 | 0.2 | 12.445 | -9.99 | -1.1 | 9.008 | -9.99 | -0.3 |
| PSO | J002.1073-06.4345 | 2.10739 | -6.43456 | 16.809 | 0.107 | 10.1 | 15.684 | 0.141 | 7.7 | 11.892 | -9.99 | 1.5 | 8.759 | -9.99 | 0.2 |
| SDWISE | J0008+3616 | 2.2142917 | 36.2704138 | 16.045 | 0.052 | 20.7 | 15.373 | 0.092 | 11.8 | 12.043 | -9.99 | 1.8 | 8.786 | -9.99 | 1.1 |
| PSO | J002.3786+32.8702 | 2.37870183 | 32.87026179 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -9.99 | -9.99 | -9.9 |
| SDSS | J0017-1000 | 4.3111666 | -10.01539722 | 15.936 | 0.055 | 19.7 | 15.167 | 0.094 | 11.5 | 12.026 | 0.334 | 3.2 | 8.52 | -9.99 | 1.2 |
| PSO | J004.3936+17.0862 | 4.39361347 | 17.08630447 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -99.99 | -9.99 | -9.9 | -9.99 | -9.99 | -9.9 |
| PSO | J004.8140-24.2991 | 4.81408 | -24.29916 | 16.281 | 0.069 | 15.8 | 15.569 | 0.116 | 9.4 | 12.123 | 0.344 | 3.2 | 8.82 | -9.99 | 0.5 |
| VDES | J0020-3653 | 5.1311237 | -36.8949476 | 16.844 | 0.094 | 11.6 | 16.354 | 0.204 | 5.3 | 12.679 | -9.99 | -0.1 | 8.342 | -9.99 | 0.8 |

Table 3. The mid-infrared photometric properties from the WISE ALLWISE catalogue for the 424 very-high redshift quasars. The first ten objects are given here as guidance to the format of the data table. The full table can be found online. *This is the third table here; the SECOND table is this with the NIR data...*

2.1 Spectroscopy

We have obtained a list of 424 spectroscopically confirmed quasars with redshifts $z \geq 5.00$.

In Table 1 we give the discovery reference for the VHzQs noting that some objects were discovered independently and contemporaneously. The redshifts for the VHzQs generally come from the measurement of broad UV/optical emission lines. Where there are far infra-red emission lines e.g. C II 158 micron, we report these, but at the level of our current analysis broadband redshifts are sufficient.

Specifically, we use data from: Fan et al. (2000), Fan et al. (2001), Fan et al. (2003), Fan et al. (2004), Mahabal et al. (2005), Cool et al. (2006), Fan et al. (2006), Goto (2006), McGreer et al. (2006), Carilli et al. (2007), Kurk et al. (2007), Stern et al. (2007), Venemans et al. (2007), Willott et al. (2007), Jiang et al. (2008), Wang et al. (2008), Jiang et al. (2009), Kurk et al. (2009), Mortlock et al. (2009), Willott et al. (2009), Carilli et al. (2010), Wang et al. (2010), Willott et al. (2010), Willott et al. (2010), De Rosa et al. (2011), Mortlock et al. (2011), Wang et al. (2011), Zeimann et al. (2011), Morganson et al. (2012), Venemans et al. (2012), McGreer et al. (2013), Venemans et al. (2013), Wang et al. (2013), Willott et al. (2013), Bañados et al. (2014), Calura et al. (2014), Leipski et al. (2014), Bañados et al. (2015), Bañados et al. (2015), Becker et al. (2015), Carnall et al. (2015), Jiang et al. (2015), Kashikawa et al. (2015), Kim et al. (2015), Reed et al. (2015), Venemans et al. (2015b), Venemans et al. (2015a), Willott et al. (2015), Wu et al. (2015), Venemans et al. (2016), Wang et al. (2016), Matsuoka et al. (2016), Wang et al. (2016b), Mortlock et al. (2011), McGreer et al. (2013), Venemans et al. (2013), Venemans et al. (2013), Venemans et al. (2015b), Venemans et al. (2015a), Bañados et al. (2016), Matsuoka et al. (2016), Reed et al. (2017), Wang et al. (2017), Mazzucchelli et al. (2017), Ikeda et al. (2017), Tang et al. (2017), Koptelova et al. (2017), Bañados et al. (2018), Matsuoka et al. (2018a) and Matsuoka et al. (2018b).

Table 1 gives the salient details for the objects used in this study. We use all the $z \geq 5.00$ quasars that have been discovered and spectroscopically confirmed as of the time of writing (2018 June). We report near-infrared ($yYJHK$ -bands) and mid-infrared (WISE W1/2/3/4) photometry and give our calculated M_{1450} .

2.2 Optical Photometry

2.2.1 Pan-STARRS1 (PS1)

We query the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS)⁷ Data Release 1 (DR1) Catalog Archive Server Jobs System (CasJobs) service at [mas-tweb.stsci.edu/ps1casjobs/](https://outerspace.stsci.edu/display/PANSTARRS). The PS1 survey observed the 30,000 deg² of sky north of declination -30 degrees in five filters *grizy*. Pan-STARRS1 (PS1) is the first part of Pan-STARRS to be completed and is the basis for the DR1. Chambers et al. (2016), Magnier et al. (2016c), Waters et al. (2016), Magnier et al. (2016a), Magnier et al. (2016b) and Flewelling et al. (2016) describe the instrument, survey, and

| Survey | # VHzQs | Notes/Survey reference |
|---------|-----------------|----------------------------|
| ATLAS | 4 | Shanks et al. (2015) |
| CFHQS | 20 | Willott et al. (2007) |
| DELS | 2 | Dey et al. (2018) |
| ELAIS | 1 | Väisänen et al. (2000) |
| FIRST | 1 | Becker et al. (1995) |
| IMS | 1 | Kim et al. (2015) |
| MMT | 12 | McGreer et al. (2013) |
| NDWFS | 1 | Jannuzi & Dey (1999) |
| PSO | 84 | Kaiser et al. (2002, 2010) |
| RD | 1 | Mahabal et al. (2005) |
| SDSS | 156 | Stoughton et al. (2002) |
| SDUV | 20 | Yang et al. (2017) |
| SDWISE | 27 | Wang et al. (2016a) |
| SHELLQs | ^a 63 | Matsuoka et al. (2016) |
| ULAS | 10 | Lawrence et al. (2007) |
| VDES | 11 | Reed et al. (2017) |
| VIK | 9 | Edge et al. (2013) |
| VIMOS | 1 | Le Fèvre et al. (2003) |

Table 4. The number of VHzQ from given surveys, with the key survey or telescope reference. ^aIncludes 8 objects with a Hyper-SuprimeCam (HSC; Miyazaki et al. 2018) designation.

data analyses. The principal science product of the PS1 survey is the catalog accessible through the CasJobs interface.

We query and return the mean PSF magnitudes from the *grizy* filters (**MeanPSFMag**) which are in the AB system for our 424 VHzQ sample. Details of our SQL and links to the main tables are given in Appendix C.

2.3 Near-infrared photometry

Due to their selection of being very faint/undetected in the observed-frame optical but bright in the observed frame near-infrared, VHzQs are generally detected in the near-infrared $yYJHK$ -bands (≈ 0.98 – $2.38 \mu\text{m}$; e.g., Peth et al. 2011).

The near-infrared data in this paper comes from the Wide Field Astronomy Unit's (WFAU) Science Archives for UKIRT-WFCAM, the WFCAM Science Archive (WSA Hambly et al. 2008) and VISTA-VIRCAM, the VISTA Science Archive (VSA Cross et al. 2012). These archives were developed for the VISTA Data Flow System (VDFS Emerson et al. 2004).

We access both the WSA and the VSA and include all non-proprietary WFCAM data, which covers all public surveys and PI projects from Semester 05A to 1st January 2017, and all non-proprietary VISTA data, which covers all public surveys and PI projects from science verification (20091015) to 1st April 2016.

The data was processed using a matched-aperture photometry method where flux is recovered at the spectroscopic position of the quasar, without necessarily knowing if there is a formal detection in the NIR photometry beforehand. Full details of the matched-aperture pipeline will appear in a forthcoming paper, Cross et al. 2018, in prep, and has also been discussed in Cross et al. (2013).

We query the WSA and VSA performing matched-aperture photometry at the positions of our 424 VHzQs. We return a database with *xxxx* entries. This database is world-readable and we give the full receipt and relevant SQL queries for accessing both databases in Appendix E.

⁷ <https://outerspace.stsci.edu/display/PANSTARRS>

2.3.1 Averaging matched photometry

The photometry in a single epoch image often has low signal-to-noise. The advantage of matched aperture photometry on QSOs is that co-adding is relatively simple if each epoch is taken in the same aperture and the aperture photometry has been corrected to total. Indeed, the standard aperture corrections work well for point sources. Coadding using the matched-aperture photometry, where the individual epochs are taken from multiple projects with different pointings and orientations, should help with issues such as scattered light, pixel distortion and aperture corrections.

We average the aperture corrected calibrated fluxes (e.g. **aperJky3**), and then convert to magnitudes. Since we do not have a deep image for each set of averages, we cannot calculate non-aperture corrected values, so the photometry is only appropriate for point-sources.

$$\bar{F} = \frac{\sum_i^N (w_i F_i)}{\sum_i^N w_i} \quad (1)$$

where F_i is the i^{th} epoch measurement of a parameter to be averaged such as the aperture corrected calibrated flux in a $1''$ aperture (**aperJky3**) and \bar{F} is the weighted mean average of this parameter. The weight for each epoch $w_i = 1/(\sigma_F)^2$ if the epoch is included and $w_i = 0$ if an epoch is excluded for quality control purposes.

We calculate a set of averaged catalogues, for each pointing and filter, based on the requirements in **RequiredMapAverages**, in these cases over time spans of 7, 14, 30, 91, 183 days, 365 days, 730 days, over 10 epochs and over all epochs, see Tab. ???. The averaging process starts at the first epoch and works on.

We detect 304 unique quasars in the WFCAM WSA database, 203 quasars are detected in the VISTA VSA database with 114 objects in common with both WFCAM and VISTA data. We give the necessary SQL queries syntax in Appendix E.

2.4 MIR data

The MIR data for this study comes exclusively from the Wide-field Infrared Survey Explorer (WISE) mission. Since we are only concerned here with the very large area ($\gg 1000 \text{ deg}^2$) surveys, we leave exploration of the VHzQ population in e.g. the large Spitzer areal surveys such as the Spitzer IRAC Equatorial Survey (SpIES; Timlin et al. 2016), the Spitzer-HETDEX Exploratory Large-area Survey (SHELA; Papovich et al. 2016) and the Spitzer-SPT Deep Field (SSDF; Ashby et al. 2013), to a future investigation. This will also include a detailed study of MIR spectra e.g. Lambrides et al. (2018).

We use data from the the beginning of the WISE mission (2010 January; Wright et al. 2010) through the fourth-year of NEOWISE-R operations (Mainzer et al. 2011, 2017 December;). More specifically, we use the data from the **AllWISE** program and catalogue, which combines data from the WISE cryogenic and NEOWISE (Mainzer et al. 2011 ApJ, 731, 53) post-cryogenic survey phases. For the our variability investigations, we supplement the ALLWISE data with data from the **NEOWISE 2018 Data Release**. NEOWISE

| Selection | number detected (%) |
|--|---------------------|
| Any band (<i>YJHK/K_s</i>) | 394 () |
| <i>Y</i> -band | () |
| <i>J</i> -band | () |
| <i>H</i> -band | () |
| <i>K</i> or <i>K_s</i> -band | () |

Table 5. Detection rate of VHzQs in the near-infrared.

2018 makes available the 3.4 and 4.6 m (W1 and W2) single-exposure images and extracted source information that were acquired between 2016 December 13 and 2017 December 13 UTC, which was the fourth year of survey operations of the Near-Earth Object Wide-field Infrared Survey Explorer Re-activation Mission (NEOWISE; Mainzer et al. 2014, ApJ, 792, 30). The fourth year NEOWISE data products are concatenated with those from the first three years (originally released on March 26, 2015, March 23, 2016 and June 1, 2017) into a single archive.

The WISE scan pattern leads to coverage of the full-sky approximately once every six months (a “sky pass”), but the satellite was placed in hibernation in 2011 February and then reactivated in 2013 October. Hence, our light curves have a cadence of 6 months with a 32 month sampling gap.

3 RESULTS

Having collated the sample of 424 VHzQs, and obtained their optical, near- and mid-infrared photometry we report here the various photometric properities of the quasars.

First, we will concentrate on detection rate in the infrared, go on to report on the color-redshift and color-color properties of our sample and then report on how the current sample populates the luminosity-redshift Lz -plane.

3.1 Detection Rates in the NIR

Table 5 gives the detection rates for the VHzQs in the NIR *YJHK/K_s*-bands. The first thing to note is that the coverage of the NIR surveys for example from the UKIDSS LAS and VISTA VHS, does not overlap the full area for where the VHzQs are detected.

3.1.1 Comparing WFCAM and VISTA

There are 114 overlapping QSOs between WFCAM and VISTA. Using the **VegaToAB** value⁸ to put these objects on the same AB system, and for each object compared the two measurements. First, the calculated weighted average (calibrated flux) in each filter of both and calculated the ratio and difference between each measurement and the average. Then for each filter we calculated the weighted average of the differences (in mag) for each instrument to see if there were significant offsets. The results are given in Table 6. The only filter with a significant offset is the *Y*-band. All of the VISTA averages are negative and all of the WFCAM ones are positive. The *K_s* versus *K* band may be slightly dodgy, given the different shapes of the filters.

⁸ What is this exactly??

| abs(VIRCAM - WFCAM) | millimag | no. of objects |
|------------------------|----------|-------------------|
| Z | 19.3 | 2 |
| Y | 66.2 | 48 |
| J | 3.2 | 105 |
| H | 19.3 | 89 |
| K_s/K | 12.7 | 93 |

Table 6. Comparing the magnitudes in different WFCAM/UKIRT and VIRCAM/VISTA near-infrared bands.

| Selection | number detected (%) |
|--------------------------------------|---------------------|
| W1 SNR > 2.0 | 275 (64.9) |
| W2 SNR > 2.0 | 255 (60.1) |
| W1 \wedge W2 SNR > 2.0 | |
| W3 SNR > 2.0 | 99 (23.3) |
| W4 SNR > 2.0 | 29 (6.8) |
| Any W1/2/3/4 SNR > 2.0 | |
| W1/2 SNR < 2.0 \wedge W3 SNR > 2.0 | |

Table 7. ATLAS Shanks et al. (2015);

3.2 Detection Rates in the MIR

Unlike the NIR coverage, the WISE satellite and mission performed an all-sky survey, so the location of every VH z Q in our dataset is covered. However, the depth of the WISE ALLWISE survey depends heavily on sky location, with locations near the Ecliptic Poles having the highest number of exposures.

Before reporting on the detection rates, we investigate this effect. Figure 3 shows the WISE magnitude versus signal-to-noise, colour coded by $wxcov$ the mean coverage depth, in each corresponding band. In the two shorter bands W1/2 we see the clear and expected trend for brighter objects to have larger SNR, and also for the higher signal to noise for objects with more exposures at a given magnitude. The behaviour for the W3/4 bands is different, with two populations clearly evident in W3 and although a bit more mixed, also in W4. With the suggested split at SNR > 2, and no obvious R.A./Declination dependence seen, this behaviour is explained by the fact that there are non-detections in W3/4 for objects (with high W1/2 SNR) that are reported in the ALLWISE catalogue.

For the 278 VH z Q with coverage detections, the mean number of exposures for the W1/2 bands is 32.0 and 31.5, respectively, with a minimum number of exposures 17 and 12, and the maximum number of exposures being 114 (for both bands). For the W3/4 filters, the corresponding mean, minimum and maximum exposure are 17.4 and 17.5, 5.8 and 6.8 and 69 (for both bands). These values are directly from the $wxcov$ entries in the WISE ALLWISE catalogue.

Table 8 gives the detection rates for the VH z Qs in the MIR WISE W1-4 bands.

Blain et al. (2013)

Recently, Assef et al. (2018) released two large catalogues of AGN candidates identified across 30,000 deg² of extragalactic sky from the WISE AllWISE Data Release. The “R90” catalogue, is contains 4.5M AGN candidates at 90% reliability (and ≈ 150 AGN candidates per deg²) while the “C75” catalog consists of 20.9M AGN candidates at 75% completeness (and ≈ 700 AGN candidates per deg²). Cross-matching out catalogue of 424 VH z Qs with these catalogues,

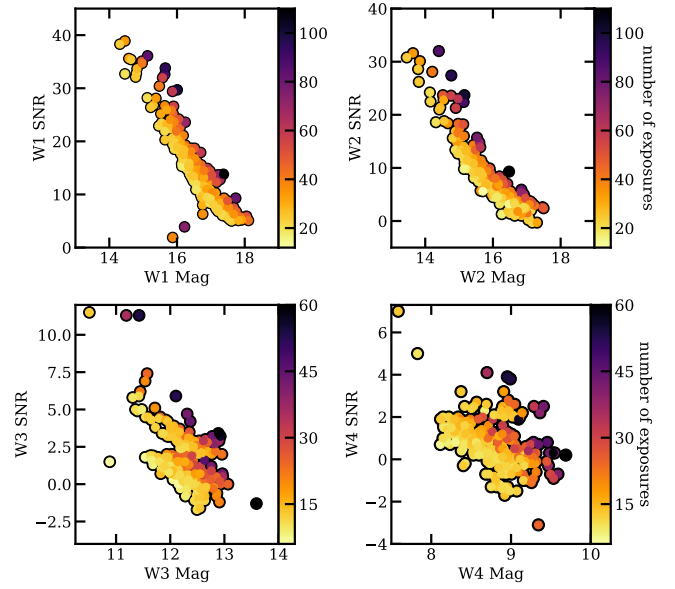


Figure 2. WISE W1/2/3/4 magnitude against signal-to-noise, colour coded by $wxcov$ the mean coverage depth, in each corresponding band.

| Selection | number detected (% of full spectra) |
|--------------------------------------|--|
| From “Source”, “Rejects”, | 245, 40 (67.2) |
| W1 SNR > 2.0 | 279 (65.8) |
| W2 SNR > 2.0 | 258 (60.8) |
| W1 \wedge W2 SNR > 2.0 | 253 (59.7) |
| W3 SNR > 2.0 | 97 (22.9) |
| W4 SNR > 2.0 | 33 (7.8) |
| W1/2 SNR < 2.0 \wedge W3 SNR > 2.0 | 3 (0.7) |

Table 8. Data from the AllWISE Source Catalog and AllWISE Reject Table, from the NASA/IPAC Infrared Science Archive

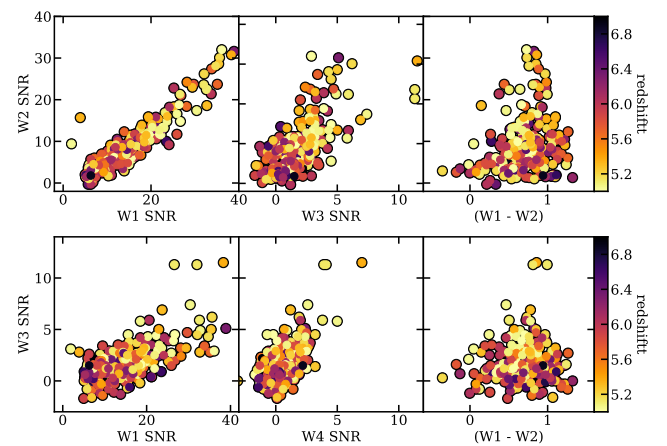


Figure 3. WISE signal-to-noise measures for the four bands, as well as for (W1-W2) colour. The points are colour coded by redshift.

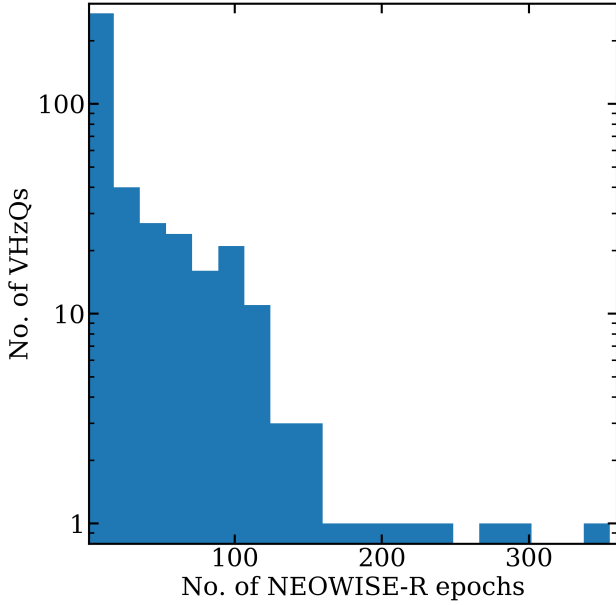


Figure 4. Histogram showing the number of NEOWISE-R epochs and detections there are for each VHzQ.

produces 42 matches with the R90 sample and 98 matches with the C75 sample. Both catalogues unsurprisingly match to the ultraluminous quasar SDSS J0100+2802 (Wu et al. 2015) while the C75, but not the R90 catalogue matches to ULAS J1120+0641 (Mortlock et al. 2011). Neither catalogue matches J1342+0928 (Bañados et al. 2018).

Very High- z Quasars Detected in WISE W3 and W4.

3.3 Variability

VHzQs, if accreting at, or above the Eddington Limit, might well have large values of changing mass accretion rate, \dot{m}_{accr} . A consequence of this would be that these quasar exhibit signs of variability, most likely showing up in their UV/optical rest-frame spectra. We look for evidence of this variability signature in the NIR and MIR light-curves of the VHzQs. As a guide, C IV enters the Y -band at redshift $z=5.32$ and exits at $z=5.99$, and enters the J -band at redshift $z=6.55$ and exits at $z=7.57$. Mg II enters the H -band at redshift $z=4.33$ and exits at $z=5.37$ and enters the K -band at redshift $z=6.25$ and exits at $z=7.50$.

Using the extended datasets described in Section 2.3 and ??, we

Figure 4 gives the number of NEOWISE-R epochs and detections there are for each VHzQ, while Figure ?? presents three examples of the MIR lightcurves and associated colour changes. Here we show J0100+2802 (Wu et al. 2015), J0224-4711 and J1626+2751.

3.4 Colours

Currently, very high-redshift quasars are identified by their morphology, flux and colours in optical and infrared imaging data Fan (1999); ? Quasars are generally selected to be point sources, but be outliers from the stellar locus in colour

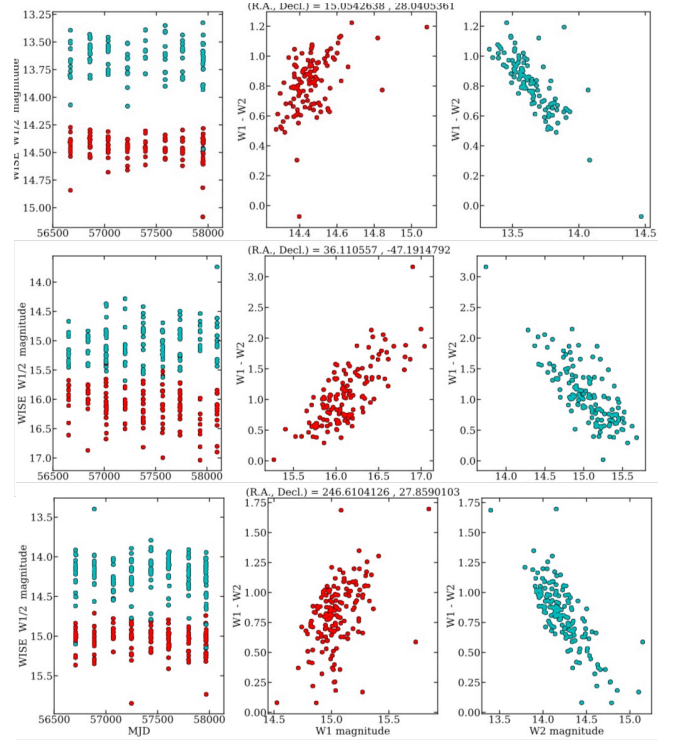


Figure 5. Here we show the MIR NEOWISE-R for J0100+2802 (Wu et al. 2015), J0224-4711 and J1626+2751. Red points are the W1 band; cyan points the W2 band.

space. For VHzQs, the main technique is to look for objects with extreme optical-to-near-infrared colours. The lack of proper motion can also help identified quasars (e.g. Lang et al. 2009).

Pellentesque vel elit neque, in interdum lacus. Quisque sodales, nunc et luctus convallis, nisl dui luctus dui, at congue urna velit a nisl. Ut sit amet sapien a risus dapibus sagittis. Cras sed ultricies erat. Donec id metus sed urna lacinia convallis vel sed enim. Proin nisi libero, ornare vel bibendum eu, sollicitudin sed leo. Cras tincidunt aliquet ultricies. Cras pretium velit leo, in malesuada enim. Duis sagittis ultricies interdum. Proin sit amet sem nec metus feugiat pharetra.

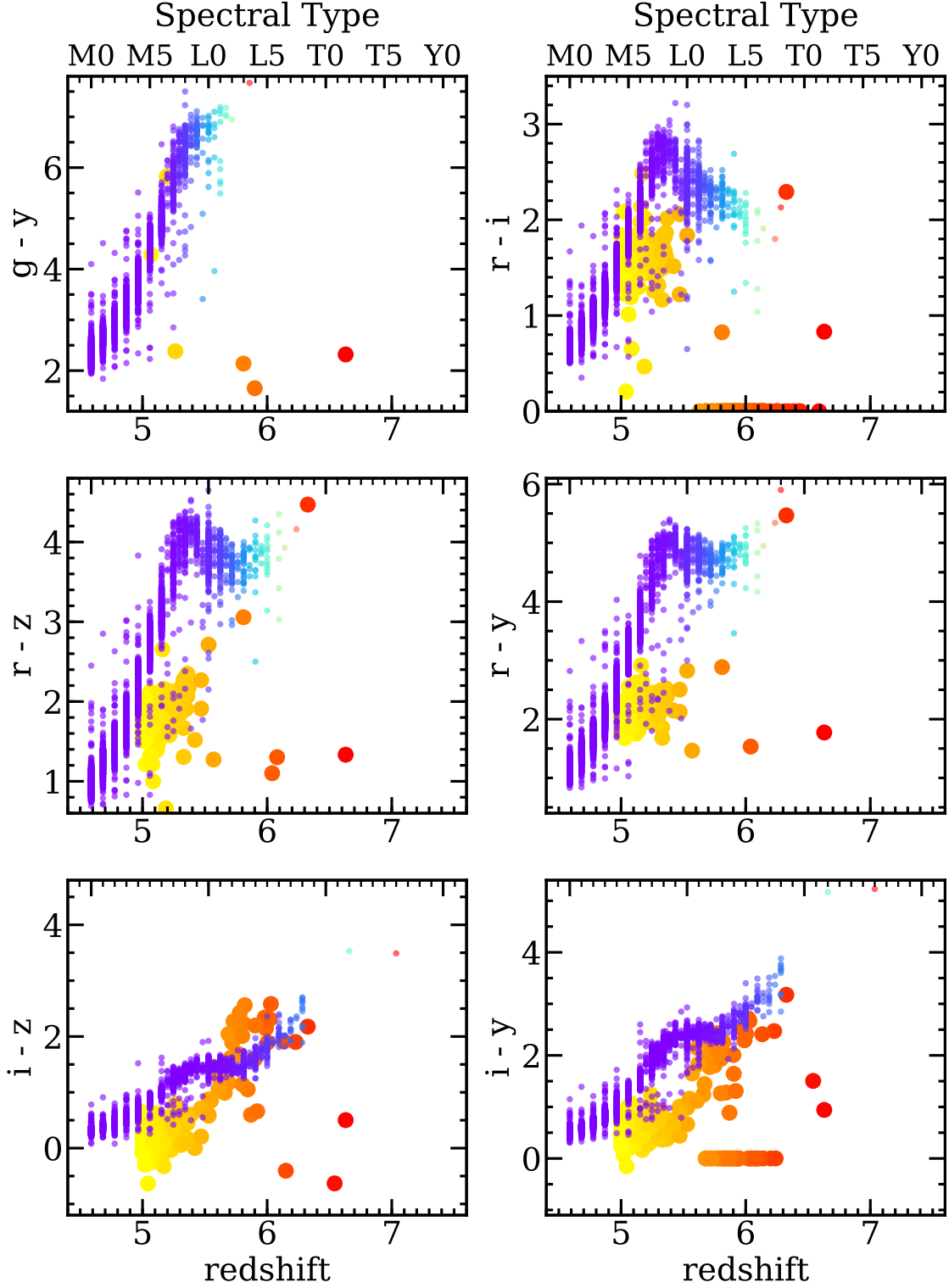
Figure 6 presents the optical colour-redshift trends for Late Type M/L/T dwarfs and the VHzQs.

Figure 6 presents the near-infrared colour-redshift trends for Late Type M/L/T dwarfs and the VHzQs.

3.5 L-z Plane

Having obtained an as-near-to-homogenous set of photometry as we can, we are now in a position to calculate the Absolute Magnitudes of the VHzQ sample and in particular the absolute magnitude at rest-frame 1450\AA , M_{1450} , which is a key physical quantity and goes directly towards the quasar luminosity function and thus the reionization of hydrogen calculation.

At $z=5.00$, the rest-frame 1450\AA emission is redshifted to 8700\AA observed, i.e., in the z -band, while at



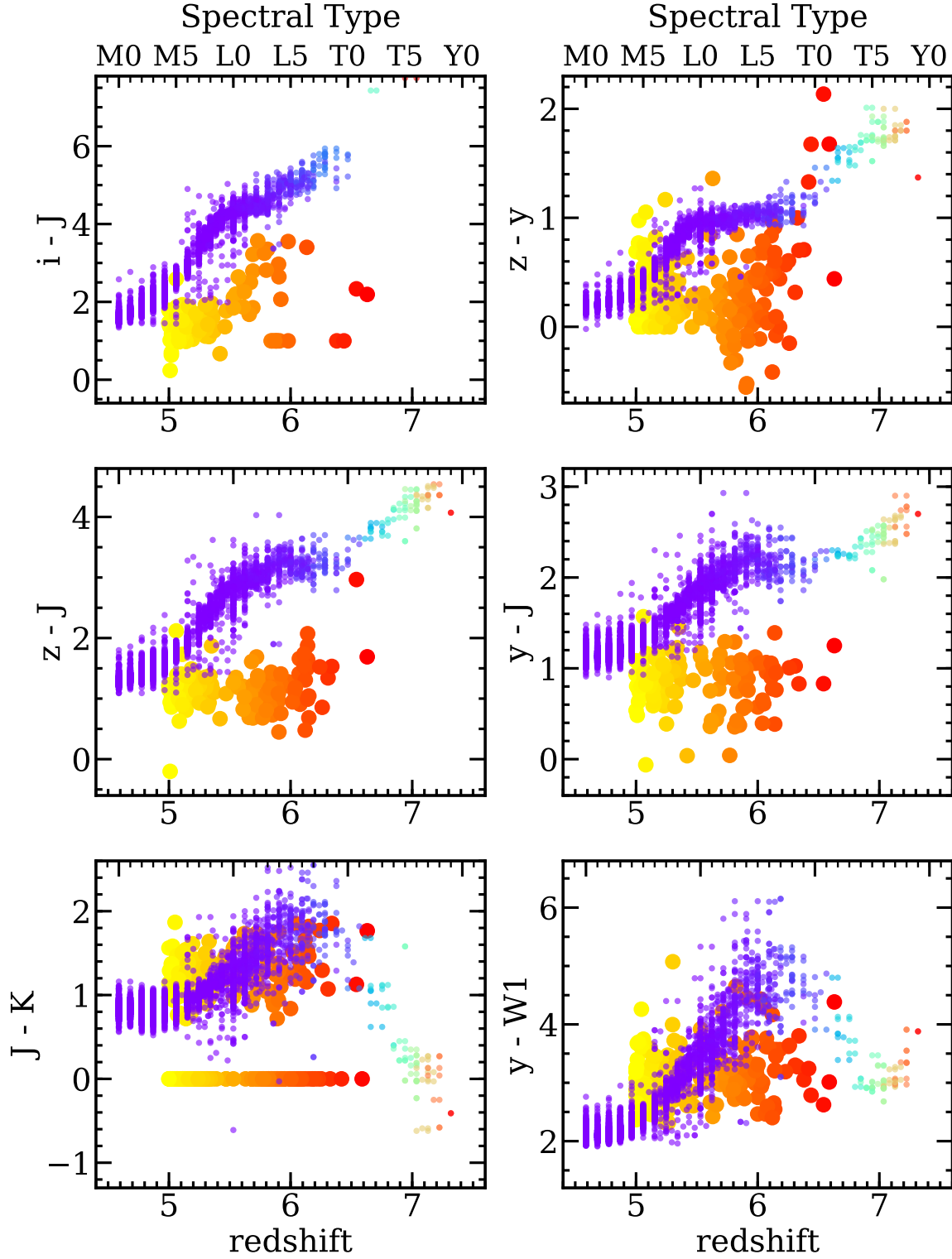


Figure 7. Infrared colour-spectral type and redshift plots for Late Type M/L/T dwarfs and the VHzQs. ^{MNRAS 000, 000–000 (0000)} *NB* I'm really not sure how Best et al. actually get their stellar sequence so clean. There are two types of spectral classification, but restricting it to just SpT_optn or SpT_nir removes the blue or red end respectively. Hmmm....

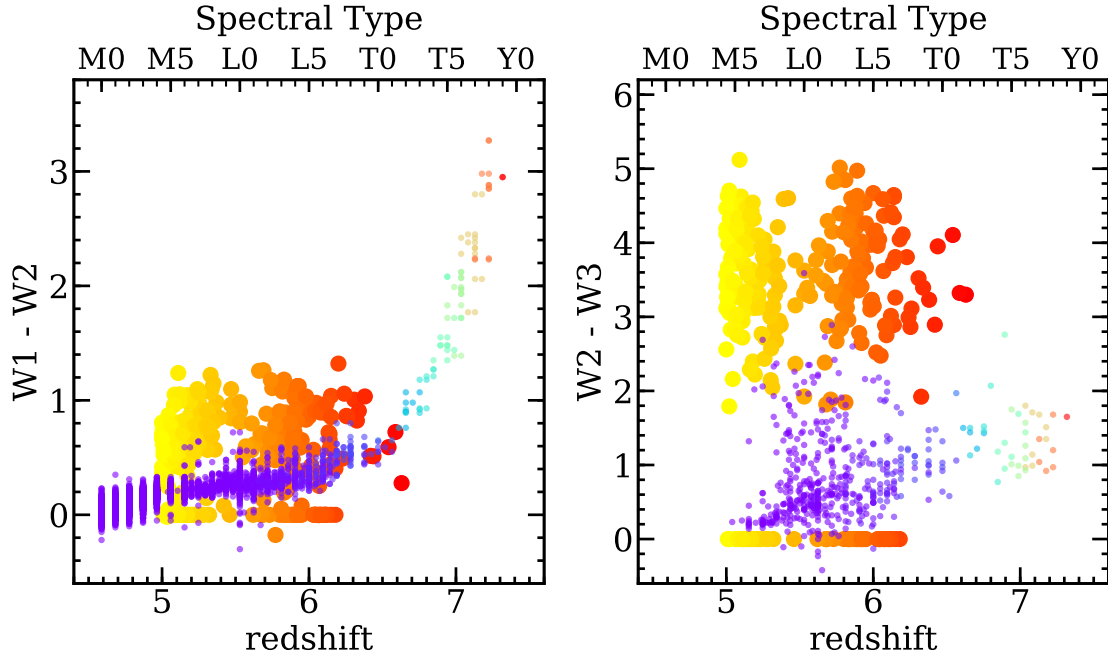


Figure 8. Infrared colour-spectral type and redshift plots for Late Type M/L/T dwarfs and the VHzQs.

3.6 SEDs and Dust properties of the VHzQs

There are a range of IR SEDs e.g. [Mullaney et al. \(2013\)](#) etc. etc. etc. However, they are, for our purposes all roughly the same.

4 DISCUSSION AND CONCLUSIONS

In this study, we have, for the first time, compiled the list of all $z > 5$ spectroscopically confirmed quasars. We have assemble the NIR ($y/Y, J, H, K/K_s$) and MIR (WISE W1/2/3/4) photometry for these objects, given their detection rates and SEDs. We find that:

We can gain a good appreciation for what these missions will discover by collating the datasets we currently have.

- Lorem ipsum dolor sit amet, consectetur adipiscing elit. Aliquam porta sodales est, vel cursus risus porta non. Vivamus vel pretium velit. Sed fringilla suscipit felis, nec iaculis lacus convallis ac.

- Fusce pellentesque condimentum dolor, quis vehicula tortor hendrerit sed. Class aptent taciti sociosqu ad litora torquent per conubia nostra, per inceptos himenaeos. Etiam interdum tristique diam eu blandit. Donec in lacinia libero.

- Sed elit massa, eleifend non sodales a, commodo ut felis. Sed id pretium velit. Vestibulum et turpis vitae quam aliquam convallis. Sed id ligula eu nulla ultrices tempus. Phasellus mattis erat quis metus dignissim malesuada. Nulla tincidunt quam volutpat nibh facilisis euismod. Cras vel auctor neque. Nam quis diam risus.

Nunc lacus nibh, convallis ac lobortis ut, tempus ac lectus. Maecenas eu elit massa. Nulla vel lacus lorem. Proin et lobortis tortor. Phasellus ultrices nisl non enim porttitor dictum. Curabitur nec nunc ac nibh ornare elementum. Nunc

ultrices hendrerit ultricies. Aliquam dapibus semper est et gravida. Etiam cursus, massa eget tempor elementum, lectus urna feugiat nisi, eget sagittis.

Author Contributions

N.P.R. initiated the project, compiled the list of $z > 5.00$ quasars, wrote most of the analysis code, developed the the plotting scripts, and developed and wrote the initial and subsequent drafts of the manuscript.

N.J.G.C. supplied the critical near-infrared expertise and database for which the bulk of the project relies. N.J.G.C. also contributed directly to the writing of the manuscript.

Availability of Data and computer analysis codes

All materials, data, code and analysis algorithms are fully available at: <https://github.com/d80b2t/VHzQ>

ACKNOWLEDGEMENTS

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We thank Mike Read at the ROE WFAU for help with the WFCAM Science Archiv (WSA), and also the VISTA Science Archive (VSA). We thank Bernie Shiao at STScI for help with the Pan-STARRS1 DR1 CasJobs interface.

This paper heavily used TOPCAT (v4.4) ([Taylor 2005, 2011](#)). This research made use of [Astropy](#), a community-developed core Python package for Astronomy ([Astropy Collaboration et al. 2013](#); [The Astropy Collaboration et al. 2018](#)).

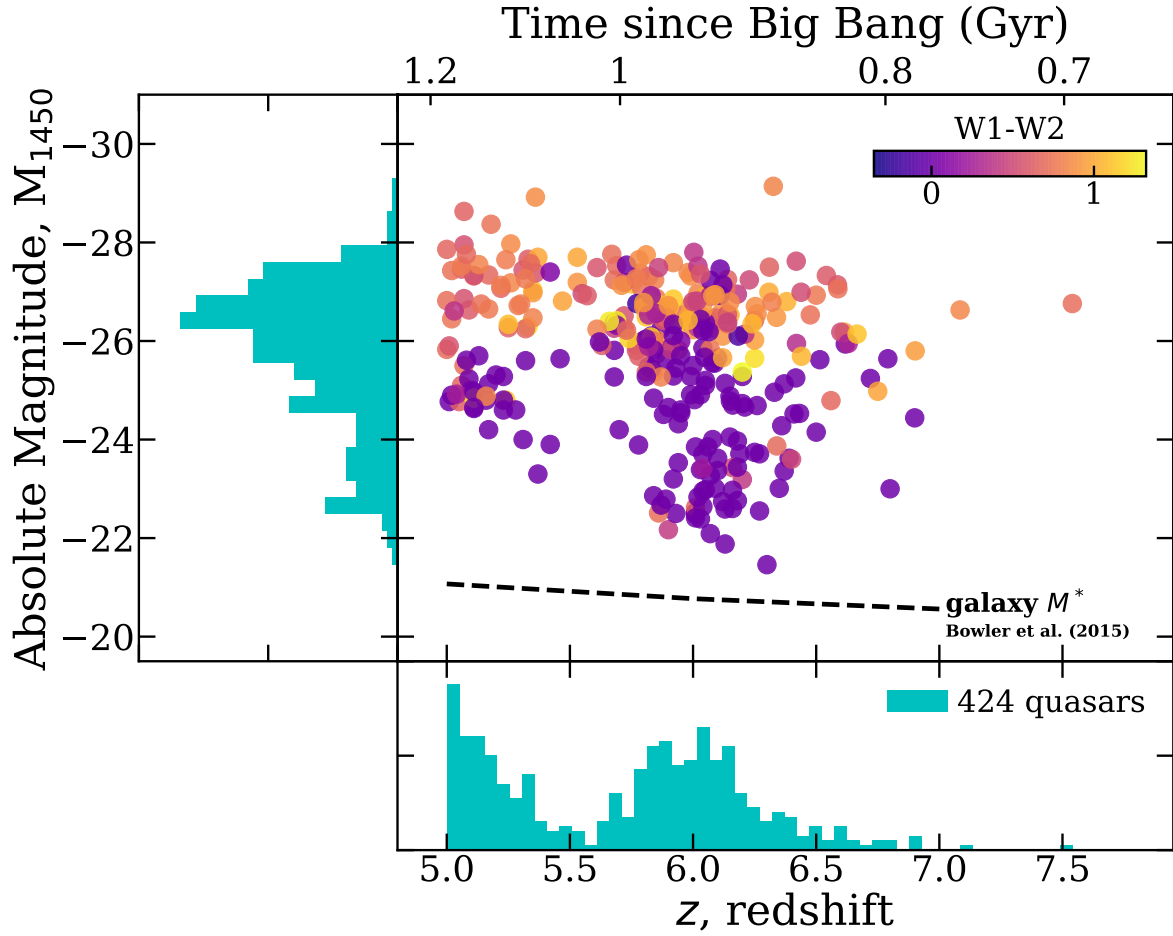


Figure 9. The spectral bands used by different survey telescopes and that are relevant here.

The Pan-STARRS1 Surveys (PS1) and the PS1 public science archive have been made possible through contributions by the Institute for Astronomy, the University of Hawaii, the Pan-STARRS Project Office, the Max-Planck Society and its participating institutes, the Max Planck Institute for Astronomy, Heidelberg and the Max Planck Institute for Extraterrestrial Physics, Garching, The Johns Hopkins University, Durham University, the University of Edinburgh, the Queen's University Belfast, the Harvard-Smithsonian Center for Astrophysics, the Las Cumbres Observatory Global Telescope Network Incorporated, the National Central University of Taiwan, the Space Telescope Science Institute, the National Aeronautics and Space Administration under Grant No. NNX08AR22G issued through the Planetary Science Division of the NASA Science Mission Directorate, the National Science Foundation Grant No. AST-1238877, the University of Maryland, Eotvos Lorand University (ELTE), the Los Alamos National Laboratory, and the Gordon and Betty Moore Foundation.

This project used data obtained with the Dark Energy Camera (DECam) and the NOAO Data Lab. The Data Lab is operated by the National Optical Astronomy Observatory,

the national center for ground-based nighttime astronomy in the United States operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, and NEOWISE, which is a project of the Jet Propulsion Laboratory/California Institute of Technology. WISE and NEOWISE are funded by the National Aeronautics and Space Administration.

CasJobs was originally developed by the Johns Hopkins University/ Sloan Digital Sky Survey (JHU/SDSS) team. With their permission, MAST used version 3.5.16 to construct CasJobs-based tools for GALEX, Kepler, the Hubble Source Catalog, and PanSTARRS.

This research has made use of the SVO Filter Profile Service (<http://svo2.cab.inta-csic.es/theory/fps/>) supported from the Spanish MINECO through grant

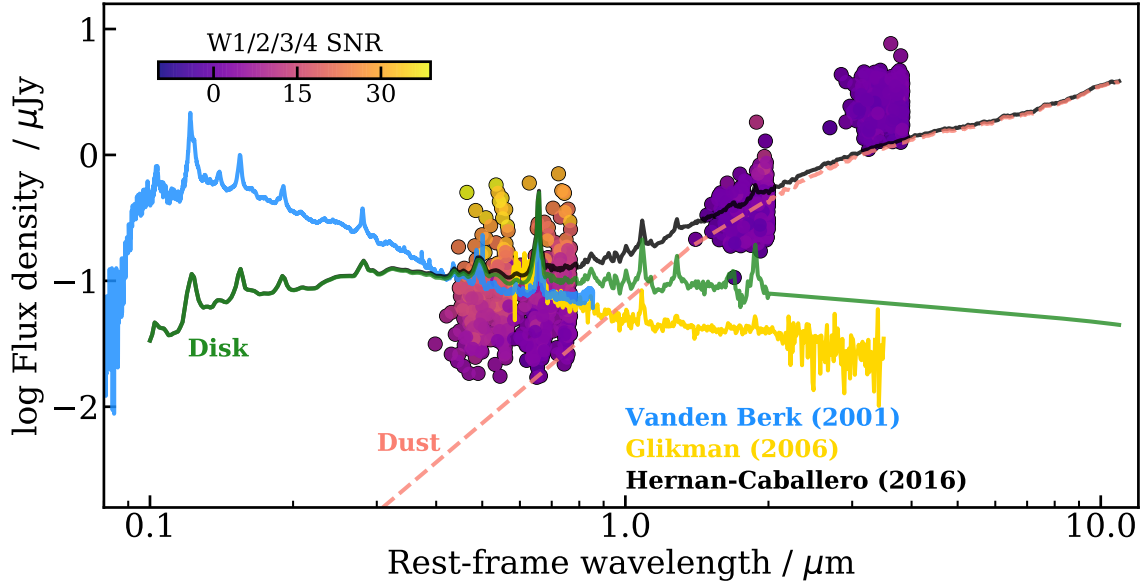


Figure 10. The rest-frame properties of the VHzQs.

AyA2014-55216 The SVO Filter Profile Service⁹ describes the Spanish VO Filter Profile Service. The Filter Profile Service Access Protocol. Rodrigo, C., Solano, E. <http://ivoa.net/documents/Notes/SVOFPSDAL/index.html>

APPENDIX A: FILTER CURVES

From the SVO Filter Profile Service¹⁰.

APPENDIX B: A. PHOTOMETRIC BANDS AND CONVERSIONS

Due to the differing normalizations between the SDSS and UKIDSS photometric systems, certain corrections are required. To present our data in the purest sense, all the NIR magnitudes from UKIDSS (originally AB magnitudes) were corrected to Vega magnitudes as suggested in Hewett et al. (2006).

Although ULAS magnitudes are reported in terms of Vega and SDSS magnitudes are reported in AB terms for the most part whenever an optical-NIR color was calculated both magnitudes were left in their default term.

<https://www.gemini.edu/sciops/instruments/magnitudes-and-fluxes>

APPENDIX C: PANSTARRS1 SQL QUERIES

The PS1 Casjobs SQL Server is located at mast-web.stsci.edu/ps1casjobs. The top level documentation is given [here](#) while the description of tables is given [here](#). The main tables are the `objectThin` and `meanObject` tables.

⁹ Rodrigo, C., Solano, E., Bayo, A. <http://ivoa.net/documents/Notes/SVOFPS/index.html>

¹⁰ <http://svo2.cab.inta-csic.es/svo/theory/fps/>

Table B1. Adapted from Table 9 of [Peth et al. \(2011\)](#). CTIO/DECam, PanSTARRS/PS1, LSST Filter only values. All wavelengths in Å. From [González-Fernández et al. \(2018\)](#) $Z_{AB} - Z_{Vega} = 0.502$; $Y_{AB} - Y_{Vega} = 0.600$; $J_{AB} - J_{Vega} = 0.916$; $H_{AB} - H_{Vega} = 1.366$; $Ks_{AB} - Ks_{Vega} = 1.827$; and the CASU Vega to AB conversions v1.3: Z,Y,J,H,Ks were: 0.524, 0.618, 0.937, 1.384, 1.839. So, Δ (vs. Gonzalez-Fernandez):: (11.2, 1.1, 5.4, 1.6, 0.1) millimags. Δ (vsCASU v1.3):: (-10.8, -16.9, -15.6, -16.4, -11.9) millimags.

| Band | λ_{eff} | λ_{min} | λ_{max} | W_{eff} | AB - Vega Transformations |
|----------------------|------------------------|------------------------|------------------------|------------------|--------------------------------------|
| g_{HSC} | 4633 | 3940 | 5546 | 1460 | $g_{\text{HSC}} = g_{AB} + 0.097$ |
| g_{LSST} | 4730 | 3877 | 5665 | 1333 | $g_{\text{LSST}} = g_{AB} + 0.083$ |
| g_{DECam} | 4734 | 3939 | 5528 | 1133 | $g_{\text{DECam}} = g_{AB} + 0.083$ |
| g_{PS1} | 4776 | 3943 | 5593 | 1167 | $g_{\text{PS1}} = g_{AB} + 0.080$ |
| r_{HSC} | 6104 | 5325 | 7071 | 1503 | $r_{\text{HSC}} = r_{AB} - 0.151$ |
| r_{PS1} | 6130 | 5386 | 7036 | 1318 | $r_{\text{PS1}} = r_{AB} - 0.153$ |
| r_{LSST} | 6139 | 5375 | 7055 | 1338 | $r_{\text{LSST}} = r_{AB} - 0.155$ |
| r_{DECam} | 6345 | 5506 | 7238 | 1379 | $r_{\text{DECam}} = r_{AB} - 0.192$ |
| i_{PS1} | 7485 | 6778 | 8304 | 1243 | $i_{\text{PS1}} = i_{AB} - 0.369$ |
| i_{LSST} | 7487 | 6765 | 8325 | 1209 | $i_{\text{LSST}} = i_{AB} - 0.369$ |
| i_{HSC} | 7633 | 6791 | 8658 | 1483 | $i_{\text{PS1}} = i_{AB} - 0.396$ |
| i_{DECam} | 7750 | 6950 | 8646 | 1371 | $i_{\text{DECam}} = i_{AB} - 0.415$ |
| z_{PS1} | 8658 | 8028 | 9346 | 966 | $z_{\text{PS1}} = z_{AB} - 0.508$ |
| z_{LSST} | 8669 | 8035 | 9375 | 994 | $z_{\text{LSST}} = z_{AB} - 0.509$ |
| Z_{VIRCAM} | 8762 | 8157 | 9400 | 978 | $Z_{\text{VIRCAM}} = z_{AB} - 0.513$ |
| Z_{WFCAM} | 8802 | 8129 | 9457 | 926 | $Z_{\text{WFCAM}} = z_{AB} - 0.514$ |
| z_{HSC} | 8915 | 8280 | 9498 | 793 | $Z_{\text{HSC}} = z_{AB} - 0.512$ |
| z_{DECam} | 9216 | 8360 | 10166 | 1502 | $z_{\text{DECam}} = z_{AB} - 0.521$ |
| y_{PS1} | 9603 | 9100 | 10838 | 615 | $y_{\text{PS1}} = y_{AB} - 0.541$ |
| y_{LSST} | 9677 | 9089 | 10859 | 810 | $y_{\text{LSST}} = y_{AB} - 0.546$ |
| Y_{DECam} | 9876 | 9355 | 10730 | 676 | $Y_{\text{DECam}} = Y_{AB} - 0.570$ |
| Y_{HSC} | 9976 | 9000 | 10931 | 1386 | $Y_{\text{HSC}} = Y_{AB} - 0.580$ |
| Y_{WFCAM} | 10305 | 9790 | 10810 | 1020 | $Y_{\text{WFCAM}} = Y_{AB} - 0.617$ |
| Y_{VIRCAM} | 10184 | 9427 | 10977 | 905 | $Y_{\text{VISTA}} = Y_{AB} - 0.601$ |
| $J_{2\text{MASS}}$ | 12350 | 10806 | 14068 | 1624 | $J_{2\text{MASS}} = J_{AB} - 0.894$ |
| J_{VIRCAM} | 12464 | 11427 | 13759 | 1628 | $J_{\text{VISTA}} = J_{AB} - 0.921$ |
| J_{WFCAM} | 12483 | 11690 | 13280 | 1590 | $J_{\text{WFCAM}} = J_{AB} - 0.919$ |
| W_{Wircam} | 14514 | 13890 | 15166 | 1020 | $W_{\text{Wircam}} = W_{AB} - 1.163$ |
| H_{WFCAM} | 16313 | 14920 | 17840 | 2920 | $H_{\text{WFCAM}} = H_{AB} - 1.379$ |
| H_{VIRCAM} | 16310 | 14604 | 18422 | 2833 | $H_{\text{VISTA}} = H_{AB} - 1.368$ |
| $H_{2\text{MASS}}$ | 16620 | 14787 | 18231 | 2509 | $H_{2\text{MASS}} = H_{AB} - 1.374$ |
| Ks_{VIRCAM} | 21337 | 19333 | 23674 | 3055 | $Ks_{\text{VISTA}} = Ks_{AB} - 1.83$ |
| $Ks_{2\text{MASS}}$ | 21590 | 19544 | 23552 | 2619 | $Ks_{2\text{MASS}} = Ks_{AB} - 1.84$ |
| K_{WFCAM} | 22010 | 20290 | 23800 | 3510 | $K_{\text{WFCAM}} = K_{AB} - 1.9$ |
| WISE W1 | 33526 | 27541 | 38724 | 6626 | $W1 = W1_{AB} - 2.699$ |
| WISE W2 | 46028 | 39633 | 53414 | 10423 | $W2 = W2_{AB} - 3.339$ |
| WISE W3 | 115608 | 74430 | 172613 | 55056 | $W3 = W3_{AB} - 5.174$ |
| WISE W4 | 228172 | 195201 | 279107 | 41017 | $W4 = W4_{AB} - 6.66$ |

```

1 SELECT s.ra, s.decl,
2        o.objID, o.raMean, o.decMean,
3        o.nDetections, o.ng, o.nr, o.ni, o.nz, o.ny,
4        m.gMeanPSFMag, m.gMeanPSFMagErr, m.gMeanPSFMagStd,
5        m.rMeanPSFMag, m.rMeanPSFMagErr, m.rMeanPSFMagStd,
6        m.iMeanPSFMag, m.iMeanPSFMagErr, m.iMeanPSFMagStd,
7        m.zMeanPSFMag, m.zMeanPSFMagErr, m.zMeanPSFMagStd,
8        m.yMeanPSFMag, m.yMeanPSFMagErr, m.yMeanPSFMagStd,
9        s.jmag, s.jmag_error, s.hmag, s.hmag_error, s.kmag, s.kmag_error into mydb.MyTable_0 from M
10

```

```

11 cross apply fGetNearbyObjEq(s.ra,s.decl,2.0/60.0) nb
12 inner join ObjectThin o on o.objid=nb.objid and o.nDetections>1
13 inner join MeanObject m on o.objid=m.objid and o.uniquePspsOBid=m.uniquePspsOBid

```

APPENDIX D: NEAR-INFRARED WFCAM SCIENCE ARCHIVE SQL QUERIES

Here we give the receipe and SQL that returned the near-infrared photometry for the VHzQs from the WFCAM Science Archive.

The data are on the WFCAM Science Archive: wsa.roe.ac.uk. Access the User Login form wsa.roe.ac.uk/login.html with these credentials::

- Username: WSERV1000
- password: highzqso
- community: nonsurvey

Then going to the [Free Form SQL Query](#) page the Database release WSERV1000v20180716 can be accessed which contains all the data we use here.

We *nota bene* a few things. First, the quantity `aperJky3` and `aperJky3Err` are found in the `wserv1000MapRemeasAver` and `wserv1000MapRemeasurement`, so care has to be taken to return unique column names (otherwise e.g. [astropy.io.fits](#) will crash). As such, we alias `aver.aperJky3` to `aperJky3Aver` and likewise for the error quantity. Aliases will be necessary in some cases anyway, because some queries can be done sensibly on multiple instances of the same table. Other times, one may join tables on quantities such as `catalogueID` or `apertureID`, where you are meaning the same thing, but aliases would again be sensible.

Second, the RA and DEC values returned by the WSA are in radians, if used directly. To return values in degrees, use a selection with an alias, e.g. `RA as RADeg` and `DEC as DECDeg`.

Then the following SQL will return the values in Table 1.

```

1  SELECT
2  qso.qsoName, qso.ra as raJ2000, qso.dec as decJ2000,
3  aver.apertureID, aver.aperJky3 as aperJky3Aver,
4  aver.aperJky3Err as aperJky3AverErr, aver.sumWeight,
5  aver.ppErrBits as ppErrBitsAver, m.mjdObs,
6  m.filterID, remeas.aperJky3,
7  remeas.aperJky3Err,
8  w.weight, remeas.ppErrBits,
9  m.project
10
11 FROM
12 finalQsoCatalogue as qso,
13 MapApertureIDshighzQsoMap as ma,
14 wserv1000MapRemeasAver as aver,
15 wserv1000MapRemeasurement as remeas,
16 MapProvenance as v,
17 wserv1000MapAverageWeights as w,
18 MapFrameStatus as mfs,
19 Multiframe as m
20
21 WHERE
22 qso.qsoID=ma.objectID and
23 ma.apertureID=aver.apertureID and
24 aver.apertureID=remeas.apertureID and
25 aver.catalogueID=v.combicatID and
26 v.avSetupID=1 and
27 v.catalogueID=remeas.catalogueID and
28 w.combicatID=v.combicatID and
29 w.catalogueID=v.catalogueID and
30 w.apertureID=aver.apertureID and
31 mfs.catalogueID=remeas.catalogueID and
32 m.multiframeID=mfs.multiframeID and
33 mfs.programmeID=10999 and
34 mfs.mapID=1
35 order by v.combicatID, m.mjdObs

```


APPENDIX E: NEAR-INFRARED VISTA SCIENCE ARCHIVE SQL QUERIES

In a very similar manner to the WSA, we give here the details on how to access the VISTA Science Archive (VSA)

At the [VSA Login](#), enter with these credentials::

- Username: `VSERV1000`
- password: `highzqso`
- community: `proprietary`

Then head to the [Freeform SQL Query](#) page where the database release to use is `VSERV1000v20180716`.

Then the following SQL will return the values in Table 1.

```

1  SELECT
2  qso.qsoName, qso.ra as raJ2000, qso.dec as decJ2000,
3  aver.apertureID, aver.aperJky3 as aperJky3Aver,
4  aver.aperJky3Err as aperJky3AverErr, aver.sumWeight,
5  aver.ppErrBits as ppErrBitsAver, m.mjdObs,
6  m.filterID, remeas.aperJky3,
7  remeas.aperJky3Err,
8  w.weight, remeas.ppErrBits,
9  m.project
10
11 FROM
12 finalQsoCatalogue as qso,
13 MapApertureIDshighzQsoMap as ma,
14 wserv1000MapRemeasAver as aver,
15 wserv1000MapRemeasurement as remeas,
16 MapProvenance as v,
17 wserv1000MapAverageWeights as w,
18 MapFrameStatus as mfs,
19 Multiframe as m
20
21 WHERE
22 qso.qsoID=ma.objectID and
23 ma.apertureID=aver.apertureID and
24 aver.apertureID=remeas.apertureID and
25 aver.catalogueID=v.combicatID and
26 v.avSetupID=1 and
27 v.catalogueID=remeas.catalogueID and
28 w.combicatID=v.combicatID and
29 w.catalogueID=v.catalogueID and
30 w.apertureID=aver.apertureID and
31 mfs.catalogueID=remeas.catalogueID and
32 m.multiframeID=mfs.multiframeID and
33 mfs.programmeID=10999 and
34 mfs.mapID=1
35 order by v.combicatID, m.mjdObs

```

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