

Observing Jovian Decametric Radio Emissions with a Software Defined Radio Telescope

David Kirwan, Dr. Alan Davy*, and John Ronan*

Waterford Institute of Technology,
Dept of Computing and Mathematics,
Cork Rd, Waterford City, Ireland
dkirwan@tssg.org
<http://www.wit.ie>

Abstract. The abstract should summarize the contents of the paper and should contain at least 70 and at most 150 words. It should be written using the *abstract* environment.

Keywords: radio astronomy, software defined radio, digital signal processing

* Supervisors

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Glossary

DAM decameter radio emissions in the 10-100m wavelengths. 5–11, 13, 14, 16, 17, 24, 25, 32–34, 48, 52

HF High Frequency usually refers to radio frequencies in the 3MHz - 30MHz. 7

IFT Io Flux Tube is a cylindrical area of space containing magnetic field lines. 5, 6, 17

IOT Internet of Things a buzzword for everything connecting to everything. 16, 35

IQ Imaginary Quadrature signal - Complex signal data and quadrature phase values in a digital signal sample. 16, 31, 48

LOFAR Low Frequency Array - radio interferometer constructed in the north of the Netherlands and across europe <https://www.astron.nl/lofar-telescope/lofar-telescope>. 28

preamp preamplifier - is an electronic amplifier that prepares a small electrical signal for further amplification or processing. 28, 32

RF Radio Frequency is a rate of oscillation which corresponds to the frequency of a radio wave. 32

MW MW - Medium Wave frequencies usually in the wavelengths 300 kHz - 3 MHz. 30–32

RSGB RSGB - The Radio Society of Great Britain <http://rsgb.org/>. 30

SDR Software Defined Radio is the creation of radio functions within software rather than hardware. 10, 13–16, 19, 24, 27, 31, 32, 48, 53

SWR Standing Wave Ratio is the ratio of the maximum peak voltage anywhere on a line to the minimum value anywhere along a line. 25, 26

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Introduction

It was discovered in 1954 by Burke and Franklin [1955] that the planet Jupiter emits radio transmissions in the decameter (DAM) range (*10-100 m wavelengths*), and the inner Jovian satellite Io appeared to have a strong control effect on these emissions [Belcher, 1987]. Jupiter's radio emissions range between 4 MHz to 40 MHz while emitting most strongly at 8 MHz [Wilkinson and Kennewell, 1994]. Due to interference from human short wave radio sources between 4-15 MHz coupled with the attenuation of these signals below 8 MHz or the refraction off Earth's ionosphere, the majority of emissions have been observed up in the 15-25 MHz range where this interference is less [Wilkinson and Kennewell, 1994]. The emission signal strength quickly diminishes above this range for ground based listening sites [Wilkinson and Kennewell, 1994].

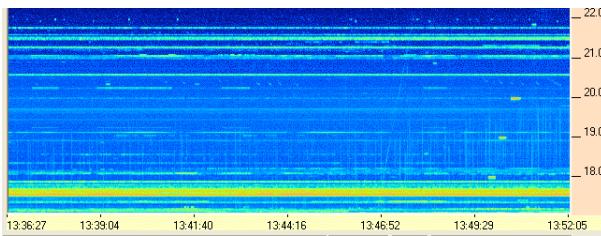


Fig. 1: Decametric Radio Emissions [Ashcraft, 2013]

Data collected by the two Voyager spacecraft in 1979 [Belcher, 1987] and the later Galileo mission in 1995 [Kivelson et al., 1996] added hugely to the understanding of the plasma interactions between Jupiter and Io and the source of the DAM emissions. It was discovered that the Io has a thin atmosphere made up of a number of neutral gasses namely sodium, potassium, sulphur, and oxygen as shown in Figure. 2. It is generally thought these gasses have been emitted through volcanic activity on the surface of the moon [Belcher, 1987]. The gasses in orbit of Io have a very short life time, due to collisions with magnetospheric electrons. This gives rise to a plasma torus (IPT) which co-rotates with Jupiter itself [Belcher, 1987]. This can also be seen in Figure. 2 which shows the IPT.

The local co-rotation speed of the plasma torus is faster than the Keplerian orbit of the moon, and the plasma overtakes Io in its orbit at 57 km/s^{-1} [Belcher,

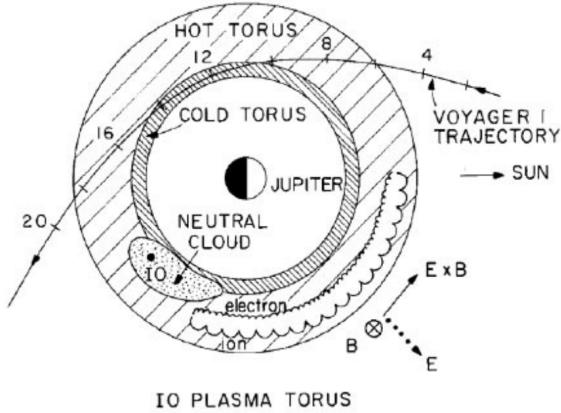


Fig. 2: Neutral Gasses in Orbit of Io [Belcher, 1987]

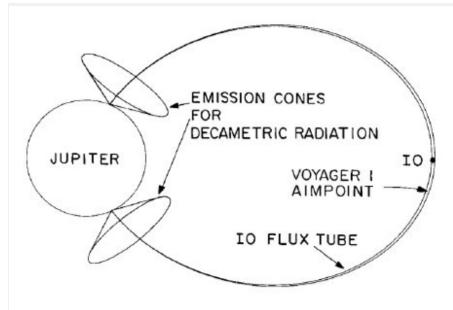


Fig. 3: Magnetic Flux Tube linking Jupiter and its satellite Io [Belcher, 1987]

1987]. Figures 3 and 40 detail diagrams of the Io Flux Tube (IFT) which is a cylinder shaped tube of space containing Jovian magnetic field lines [Belcher, 1987] which link Io to Jupiter's ionosphere at both poles. A large portion of the decametric emissions come from the area where this IFT meets the Jovian ionosphere [Belcher, 1987].

Io acts as a unipolar conductor as it orbits within this flux torus [Bose et al., 2008]. Alfvén waves are regularly produced which carry an electric charge along the magnetic field lines between Io and Jupiter [Bose et al., 2008] thereby acting as a standing wave [Bose et al., 2008]. These Alfvén waves reflect between Jupiter's ionosphere at poles and Io up to 9 times [Bose et al., 2008] before dissipating as they follow Io through its orbit.

It appears the source of the DAM emissions are largely due to these reflections of the Alfvén waves off Jupiter's ionosphere in both the northern and southern regions [Bose et al., 2008]. See Figure. 4 which shows the Alfvén wings reflecting from Jupiter's Ionosphere creating emission cones.

The DAM emissions radiate outwards in the shape of a cone as show in Figure. 3 [Belcher, 1987]. When Io is at specific points in its orbit of Jupiter, these emission cones are pointing in the direction of Earth, at which point emissions can be picked up at ground based radio telescope listening stations.

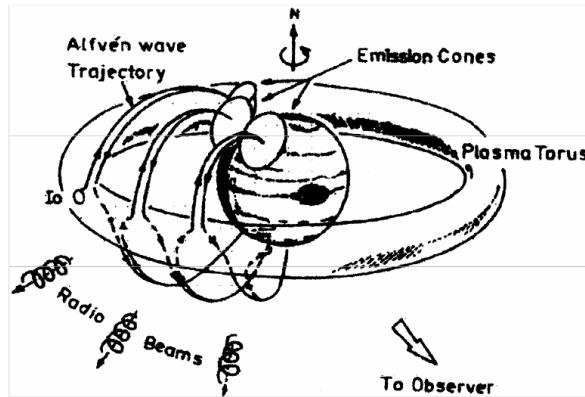


Fig. 4: Alfvén Waves following Io through its orbit and the DAM Emission cones [Bose et al., 2008]

A ground based listening station aiming to record DAM emissions from Jupiter is most likely to succeed between 15-25 MHz [Wilkinson and Kennewell, 1994]. According to ARRL [2000] there is no clear definition of the *shortwave* radio bands however it is most often considered to extend from 3 MHz to 30 MHz. ComReg is the Irish Commission for Communications Regulation within Ireland, and maintains a list of the short wave frequencies which are designated for transmission purposes in Ireland and can be seen in Figure. 5 [Comreg, 2014].

As many commercial short wave radio stations transmit in the lower end of the high frequency (HF) 3-7 MHz range, it can be extremely busy and potentially difficult to monitor DAM emissions where they are strongest. Amateur Radio operators also operate frequently in mid-late HF ranges (*160m, 80m, 60m, 40m, 30m, 20m, 17m, 15m, 12m and 10m bands*) while the higher frequency DAM emissions taper off in strength very quickly. This limits the potential listening

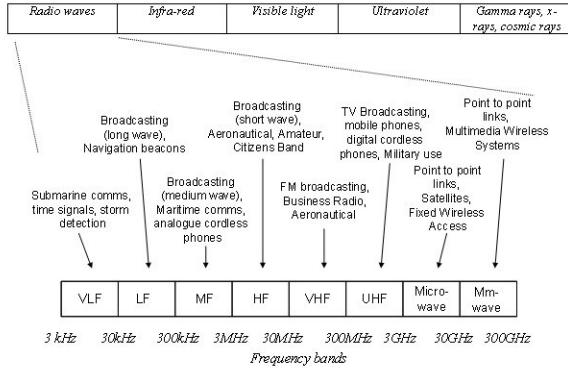


Fig. 5: Irish Regulatory Transmission Ranges [Comreg, 2014]

range significantly. Despite these obstacles, there are sections of the HF spectrum which are suitable to capture Jovian emissions. One frequency to monitor Jovian DAM emissions which is recommended by the Radio Jove project is *20.1 MHz* [NASA, 2012b].

Sourcing a suitable antenna is one of the first requirements to satisfy in order to capture DAM emissions. Antennas are generally best suited to collect electromagnetic radiation at single specific frequencies, but may resonate and therefore operate over a range of frequencies depending on the design [NASA, 2012b]. The wavelength (λ) which corresponds with the frequency (f) *20.1 MHz* can be obtained using the wavelength equation as shown in fig 6. The corresponding wavelength for the frequency *20.1 MHz* works out to be *14.925 m* using this equation.

$$\lambda = \frac{c}{f} \quad (1)$$

$$\lambda = \frac{3 \times 10^8 \text{ m/s}}{20.1 \times 10^6 \text{ Hz}} = 14.925373134328359 \text{ m} \quad (2)$$

Fig. 6: Wavelength Equation

One simple antenna design for collecting DAM emissions is the *dipole*. A dipole antenna can be constructed simply and cheaply from two pieces of wire and three insulators [NASA, 2012b], while ensuring to cut the wires to a length

matching half the desired wavelength being captured [NASA, 2012b]. However as the formula referenced in Figure. 6 describes the use of an *infinitely thin* wire which is not possible in reality, *capacitive end effects* must be taken into account when working out the resonating wavelength for a dipole antenna [NASA, 2012b].

The formula for calculating the resonating frequency for a half wavelength dipole is described in Figure: 7 and produces the value which should measure from tip to tip on the wires used to construct the dipole antenna [NASA, 2012b]. RSGB [2014] states in practice that the true resonance is not exactly multiples of the half-wavelength due to factors such as the coatings on the wire and loss due to radiation. A second formula is detailed in Figure. 8 which takes these factors into account and produces an antenna length which is slightly larger than the formula proposed by [NASA, 2012b]. Formula: 4 n stands for the number of half-wavelengths in the antenna.

$$\left(\frac{\lambda}{2}\right)m = \frac{142.65}{20.1MHz} = 7.097014925373134m \quad (3)$$

Fig. 7: Wavelength Equation for Real World Half Wavelength Dipole Antenna

$$\lambda m = 155\left(\frac{n - 0.05}{f}\right) \quad (4)$$

$$\lambda m = 155\left(\frac{1 - 0.05}{20.1MHz}\right) = 7.325870646766169m \quad (5)$$

Fig. 8: Wavelength Equation for Real World Half Wavelength Wire Antenna

The radio emissions come in several different forms each with slightly different characteristics. Table. 1 shows a list of the more widely known types which can be picked up using ground based listening equipment, and also has some information about their different characteristics [Wilkinson and Kennewell, 1994]. Any particular observation session might be made up of some or all of these different types of DAM emission and can last from a few minutes to several hours for larger noise storms [Wilkinson and Kennewell, 1994].

Type	Emission Length	Emission Description
S-Bursts	short generally 1-10 milliseconds	wideband bursts, several MHz wide
L-Bursts	long 0.5 - 5 seconds	wideband bursts, several MHz wide
N-Bursts	milliseconds upto seconds	narrowband bursts, several kHz wide

Table 1: Most common types of DAM Emissions from Jupiter [Wilkinson and Kennewell, 1994]

Figure. 9 shows an ideal case of the *S-Burst* and *L-Burst* DAM emissions and what they might look like on a frequency spectrum graph. S-Burst emissions are short, generally $1-10 \times 10^{-3}$ s long while L-Bursts can often be between 0.5-5s in length.

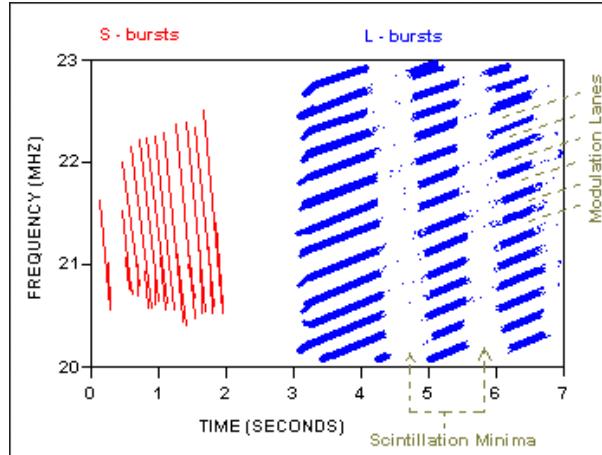


Fig. 9: Ideal DAM Emissions types from Jupiter [Wilkinson and Kennewell, 1994]

Research Hypothesis

The aim of this project is to design and construct a low cost, self sufficient software defined radio (SDR) telescope listening station, which can capture signals for transmission to a central data aggregation point for signal processing and analysis. This telescope should be suitable to study signals in the DAM (10-100 m) band at or near the *20.1 MHz* frequency in order to pick up emissions produced by either Jupiter or the Sun.

There are a number of challenges which need to be overcome to achieve this, such as Jupiter only being visible for a number of months each year and then generally in the evening, night, or morning hours. Often at highly unsociable times. For this reason a radio telescope listening site should be as automated as possible.



Fig. 10: Dual Dipole Antenna Beam at 20FT and 135deg phasing South [NASA, 2012b]

During daylight hours, the ionosphere becomes opaque to signals in the DAM band due to becoming ionised by solar radiation [NASA, 2012a]. The proposed system will have a short window in the order of 1-6 hours every second night and early morning. During which it may be possible to capture DAM emissions

from Jupiter. The Sun is a source of DAM emissions also, and the telescope can capture solar storm emissions without modification, providing the Sun passes through the antenna beam. The telescope may require manual reconfiguration in order to pick up solar storm emissions during the day.

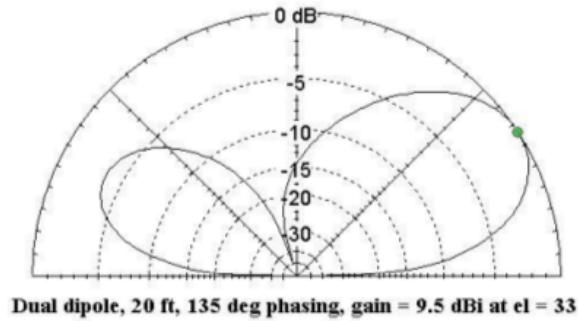


Fig. 11: Dual Dipole Antenna Array Beam[NASA, 2012b]

Interference from human sources such as short wave radio stations or amateur radio operators are also likely to affect the collection of DAM emissions from Jupiter. The ability for automated flagging or removal of interference would be a desirable feature of the system. Lightning storms can also produce interference which will affect observations, it might be desirable for the system to handle natural interference sources also.

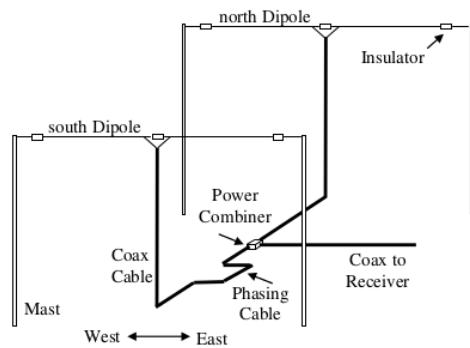


Fig. 12: Dual Dipole Antenna Array [NASA, 2012b]

Methodology

The section aims to provide a summary of the methodologies which will be followed on this project in order to answer the research questions put forward in the next section. It is broke up into the following subsections:

Antenna Build

Building the antenna is further broken down into the following steps:

- source materials to build a dual dipole antenna
- build the telescope antenna
- validate the antenna is capable of collecting signals at or near the 20.1 MHz frequency

Analysis of Listening Site Suitability

Finding a site suitable to deploy the antenna will require a site survey with the antenna connected to a spectrum analyser.

- source a spectrum analyser suitable for performing a site survey
- perform the survey at this site
- repeat until suitable site is found
- deploy the antenna at a site

Data Collection

Jupiter is visible above the horizon at night only during certain periods see Figure. 11. In order to perform the required data collection it is highly desirable to automate the process if possible.

- create listening schedule during which DAM emissions are likely to occur using the Radio Jove software
- develop SDR prototype mechanism for gathering raw data from the antenna

Data Analysis

Collected data can be analysed manually at first in order to validate the testbed, but later using software such as the development of basic algorithms to spot events such as *S-Bursts* occurring in the data.

- analyse data collected for evidence of Jovian or Solar DAM emissions
- analyse data for evidence of interference from human sources
- analyse data for evidence of interference from natural sources

- develop a system which connects to the worldwide dxcluster system and creates flag events for human identified emissions at or near the target frequency being monitored by the telescope
- develop a system which interacts with the Blitzortung servers and creates flag events for emissions generated by lightning

Data Processing

SDR data processing techniques can be developed in order to better filter the signals as they are being collected thereby minimising the level of processing at a later stage.

- using SDR techniques, to reduce interference in data as it is being collected
- attempt to remove instances of human and natural interference using signal processing techniques after collection

Data Aggregation

An API layer allows 3rd parties the ability to access data collected from multiple listening sites and potentially develop new features or systems using the data.

- develop API layer to allow external parties access aggregated data collected by the system

Design Platform

An analysis of the current available IOT technologies and some of the available options which could be used in order to produce an open low cost platform which amateur astronomers could use in order to study emissions in the DAM band.

- design a self sufficient automated platform suitable for collection of signals such as those by the SDR telescope

The SDR hardware transceiver used is likely to be the *HackRF* system while the SDR software components will be developed in either *C++* or *Python* as both languages have bindings for the *GNURadio* development tool-kit [Gnuradio, 2014]. The backend aggregation system with API will most likely be developed in the *Ruby* language as this language is highly expressive and suitable for rapid prototype development. Execution of the Ruby application on the Java VM is

capable using the *JRuby* system, which will also allow the incorporation of any required Java libraries.

Potential Pitfalls

It is difficult to determine at this early stage what will ultimately prove to be the most troublesome element to implement, however there are a number of potential pitfalls which have already been identified.



Fig. 13: Radio Jove graph showing Jupiter's max altitude 2014 [Radio-Sky Publishing, 2014]

The latitude of Ireland is 53.3° degrees N. In Ireland Jupiter reaches a maximum altitude of 52° degrees in 2014 see Figure. 13. The telescope configuration which will be documented will apply only to locations at this or a similar latitude. It will be universally applicable at this latitude without modification throughout the world, and with relatively minor modifications to the configuration of the antenna it can be used at all locations. NASA [2012b]

The SDR complexity of the solution required in order to correctly filter interference might ultimately prove too resource intensive to work with a low power system such as the *Raspberry Pi* or *Beaglebone Black*. It might become apparent that the minimum system requirements may need to be increased to an *Intel Atom* powered Netbook for example, or potentially something a lot more powerful such as an *i3* or *i5* system. If this was the case, it will drastically increase the power and cost requirements of the system should it be self sufficient, requiring bigger batteries, more capable/expensive photovoltaic and or wind turbines to keep the system topped up.

Research Questions

The initial research questions which arise are as follows:

1. *What current Internet of Things (IOT) technologies would best suit the development of a software defined radio signal listening station and how cheaply can it be created?*
 - How feasible to use wireless technologies such as Zigbee, WIFI or Bluetooth to stream collected data to a central facility?
 - If wireless technologies prove too slow, is 10/100/1000 MBit Ethernet suitable instead? Or link aggregation technologies such as EtherChannel or LACP.
 - Can a low powered computing platform such as the Raspberry Pi 2B or Beaglebone Black be used to host the system, or will a more powerful system be required?
2. *What processes or algorithms need to be developed to filter or flag known instances of human radio interference from radio signal observations?*
 - Flag transmission signals identified by amateur radio enthusiasts from a local DXSpider server in recorded data.
 - Flag instances of natural radio interference such as lightning from the Blitzortung server in recorded data.
3. *What SDR processes and algorithms need to be developed to identify instances of the three main DAM emission types detailed in Table. 1?*
 - Can existing signal demodulation techniques be employed to find DAM emissions?
 - What is required from an algorithm which might identify DAM emissions within an IQ signal?

Literature Review

The literature review can be broken down into the following areas:

- What are the decametric radio emissions and what are they caused by
- Potential radio telescope designs which could be replicated in order to collect DAM emissions
- Digital signal processing

Decametric Radio Emission what are they and where do they come from?

Belcher [1987] states that the data collected by both *Voyager* spacecraft fit with the Alfvén wing theory as an explanation for the source of the DAM emissions [Belcher, 1987]. Kivelson et al. [1996] discusses the refinements made to this theory to take into account the flowing plasma between Jupiter's ionosphere and the electrically conducting Io. The *Galileo* spacecraft collected data which appears to corroborate the updates to the Alfvén wing model [Kivelson et al., 1996]. Bose et al. [2008] states that the Alfvén waves reflect off Jupiter's ionosphere which cause the DAM emissions to radiate out from the point in Jupiter's ionosphere where it meets the IFT in a cone shape. Imai et al. [2008] proposes an update to this model to take into account the decade long shifts in DAM emission patterns, and proposes an extension to the emission cone model to include a *searchlight* shaped emission zone [Imai et al., 2008] see Figure. 39.

Radio Telescope Designs, Which to Choose?

The NASA project *Radio Jove* recommended the dual dipole array design for a cheap low cost listening site as shown in Figure. 12 [NASA, 2012b]. But it is by no means the only antenna design which would be capable of picking up the DAM emissions. Wilkinson and Kennewell [1994] recommends a slightly more advanced design, the *folded dipole* which maximises the bandwidth available to the antenna [Wilkinson and Kennewell, 1994]. Greef [2012] details a third alternative for collecting DAM emissions, a shortwave loop antenna as can be seen in Figures. 41, 42 and 43.

Digital Signal Processing

Smith [2003a] explains that digital signal processing is a term used to describe the mathematics, algorithms and techniques used to manipulate signals after they have been converted from analogue to digital. Freidt [2013] holds that digital signal processing is the preferred means to process signals, for reasons such as stability and for its resistance to long term ageing effects that analogue systems suffer from. However as Smith [2003b] points out, due to the nature of the conversion from analogue to digital, information is lost in this process. Smith [2003b] maintains that the proper sampling of an analogue signal occurred if you can reconstruct the signal from the digital samples collected, and incorrect sampling will lead to loss of the original signal due to aliasing.

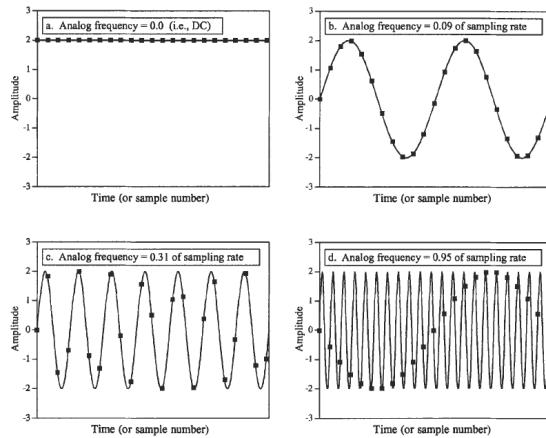


Fig. 14: Sampling an Analogue Signal [Smith, 2003b]

Smith [2003b] demonstrates in Figure: 14 how signal data loss occurs due to aliasing. When sampling an analogue signal which has a frequency more than half of the sample rate (Nyquist frequency), Smith [2003b] observes that the original signal cannot be reconstructed. This is referred to as the sampling theorem [Smith, 2003b] where f_s is the sample rate and f is the frequency being sampled.

Ossmann [2015b] contends that the application of Nyquist-Shannon sampling theorem only applies to real analogue signals where the bandwidth is equal to half the sample rate. However for a complex digital signal the bandwidth is equal to the sample rate [Ossmann, 2015b].

$$f_s > 2f \quad (6)$$

Fig. 15: Sampling Theorem aka Nyquist-Shannon sampling theorem [Smith, 2003b]

Due to the cheap availability of computational resources, digital signal processing is capable of being carried out now within software entirely, this has led to the development of software defined radio solutions [Freidt, 2013]. The GNU-Radio software allows creation of software defined radio solutions by replacing hardware functions with modular software functions. These software functions are capable of being connected together, an output from one function providing input for another. It is in this fashion it is possible to transform digital signals using digital signal processing methods within the GNURadio platform[Gnuradio, 2014].

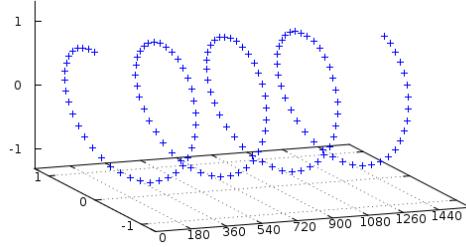


Fig. 16: I/Q Samples Helix [Kuisma, 2014]

Ossmann [2015a] illustrates how the HackRF One SDR transceiver can be accessed within GNURadio by way of the osmocomm source which is an abstraction layer between GNURadio and the HackRF driver. Figure: 46 shows an amplitude modulation signal which has been digitally sampled. In order to demodulate such a signal, an understanding how the signal data is stored is a must. Digital periodic signal samples tend to be stored in a complex format which amounts to polar coordinates [Ossmann, 2015c]. I is the point on the real axis, and Q is an angle which contains the imaginary component. See Figure: 16

shows how these I/Q values are plotted periodically and Figure: 17 shows the I/Q polar data Kuisma [2014].

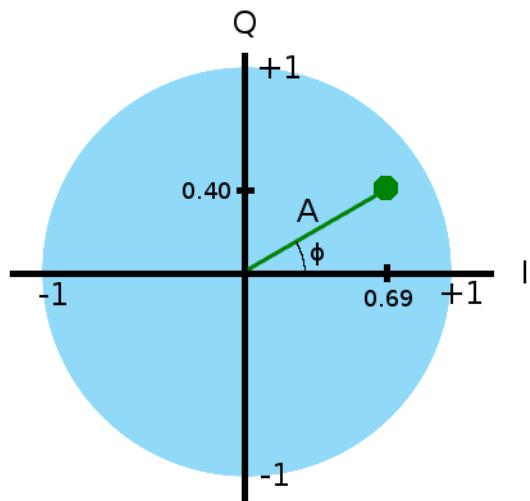


Fig. 17: I/Q Samples Polar Plot [Kuisma, 2014]

Topic	Belcher 1987	Kivelson et al 1996	Bose et al 2008	Imai et al 2008
What are the Jovian DAM emissions and what causes them?	Large fraction of DAM emissions come from point where IFT meets Jupiter's ionosphere (p3). Large moving conductor within a magnetised plasma (p1) produces Alfvén wing.	Alfvénic disturbances, generated by flowing plasma from Jovian magnetosphere and conducting Io (p337).	Alfvén waves reflecting off Jupiter's Ionosphere (p79).	Not fully understood, but believed to be produced by cyclotron maser instability [Imai et al., 2008].
Topic	Wilkinson et al 1994	NASA Radio Jove 2012	Greef 2012	
Radio telescope designs suitable to capture DAM emissions from Jupiter	Wilkinson et al suggest a folded dipole antenna design	Radio Jove project suggest a dual dipole antenna	Greef demonstrates a low cost Loop Antenna design with reflection plate	
Topic	Smith 2003	Freidt 2013	GNU Radio 2014	Ossmann 2015
Digital signal processing	Smith argues digital signal processing is one of the most powerful technologies that will shape the 21st century.	Freidt states digital signal processing now ubiquitous, and cheap availability of computational power lends self to software defined radio solutions.	GNURadio allows digital signal processing using transformation functions developed in software.	Ossmann argues that the HackRF One is perfectly suited as a low cost wide range transceiver for use within software defined radio prototypes.

Table 2: Literature Review Synthesis Matrix

Main Milestones Anticipated

The following Table. 3 details the milestones which have been identified and agreed to date. It has been updated for the Interim Report to include completed milestones and those in progress:

Deadline	Start	End	Summary
Antenna Build	October 14	March 15	Settle on a design for the telescope, source the parts for the build and finally construct the prototype antenna which will act as a template for the second dipole.
Site Survey	March 15	March 15	Perform a site survey using a spectrum analyser connected to the prototype antenna.
Deploy Dual Dipole Antenna	March 15	March 15	Deploy the antenna array at a suitable location and begin to collect data for analysis.
Interim Report	January 15	(Hard Deadline) 30th April 15	Poster presentation to review panel and supervisor on 19th May
Data Collection	March 15	July 15	Once the antenna is deployed begin collecting data for analysis.
Data Analysis and System Development	March 15	July 15	Develop analytical SDR tools to filter or flag interference. Develop algorithms to detect the various DAM emission.
Evaluate IOT Technologies	January 15	July 15	Experiment with the various IOT technologies which could potentially be used to create a self sufficient listening array.
Final Report Submission	June 15	(Hard Deadline) September 15	First draft to supervisor due in early June 15. Complete draft due 24th August. Final submission 4th September

Table 3: Milestones Anticipated

		Name	...	Start	Finish
1		Supervisor Assignment	...	08/09/14 08:00	01/10/14 17:00
2		Discuss radio telescope build with supervisors	...	07/10/14 08:00	21/10/14 17:00
3		Produce milestones	...	11/11/14 09:00	18/11/14 17:00
4		Deliver Dissertation Proposal	...	08/12/14 08:00	19/12/14 17:00
5		Dissertation Proposal Presentation	...	05/01/15 08:00	08/01/15 17:00
6		Deliver Draft Proposal to Supervisors	...	08/01/15 08:00	16/01/15 16:00
7		Deliver the Interim Report	...	03/04/15 08:00	30/04/15 17:00
8		Perform literature Review	...	08/09/14 08:00	05/06/15 17:00
9		Deliver the final thesis	...	01/09/15 08:00	01/09/15 17:00
10		Source materials used for telescope construction	...	07/10/14 08:00	03/11/14 17:00
11		Purchase materials used for the antenna pole	...	04/11/14 08:00	19/03/15 17:00
12		Construct the dipole antennas	...	14/11/14 09:00	11/12/14 17:00
13		Solar Eclipse Radio Propagation Experiment	...	20/03/15 08:00	20/03/15 17:00
14		Perform Site Suitability Survey	...	19/03/15 08:00	20/03/15 17:00
15		Investigate CoAP Wireless Protocol	...	06/05/15 08:00	08/05/15 13:00
16		Investigate WiFi Wireless Technology	...	06/05/15 08:00	08/05/15 13:00
17		Investigate Zigbee Wireless Technology	...	06/05/15 08:00	08/05/15 13:00
18		Investigate Bluetooth Wireless Technology	...	06/05/15 08:00	08/05/15 13:00
19		Investigate MQTT Wireless Protocol	...	06/05/15 08:00	08/05/15 13:00
20		Develop SDR Data Collection Prototype	...	09/02/15 08:00	01/07/15 17:00
21		Develop Spectrogram Generator	...	06/04/15 08:00	01/07/15 17:00
22		Develop Spectrogram Overlays	...	25/05/15 08:00	30/06/15 17:00

Fig. 18: SDRT Gantt Chart Details



Fig. 19: SDRT Gantt Chart

Testbed Build and Verification

The NASA Radio Jove dual dipole antenna design (Figure: 12) was chosen over the other designs for a number of reasons, namely:

- Simplicity of the architecture and consequently also construction
- Low cost of building materials
- Reasonably effective performance (5-9 dB gain depending on configuration)
- Easy to operate and maintain
- Suitable for a permanent installation
- Large body of technical documentation supporting the operation

Special Resources Required

The following resources have been identified as being required in order to perform the research and are listed as follows:

GnuRadio

The GnuRadio software is an open-source software development tool kit that allows access to signal processing techniques in order to implement software radio solutions.

HackRF

The HackRF SDR transceiver system is the software defined radio transceiver chosen for the prototype system.

PL259 / SO239

Cable adapters which connect coax cables together with the dipole center.

Radio Jove Software

The Radio Jove software is extremely useful for observers wishing to capture DAM emissions from Jupiter. It contains a large number of features such as emission prediction observation charts, and information regarding Jupiter's location in the sky from any specified point on the Earth's surface.

RG59 Coax

RG59 coax cable is required to link the dipole antennas back to the SDR transceiver.

Spectrum Analyser

A spectrum analyser will be required in order to perform a site survey to determine if a location is suitable for collecting DAM emissions.

Antenna Build

In order to satisfy the main requirement of capturing DAM signals at 20.1MHz construction of an antenna was needed. The equation shown in Figure: 8 was used to calculate a starting estimate of the length of antenna wire required to construct this antenna. Many factors can affect the ability of an antenna to receive data at a particular frequency [ARRL, 2000], such as the following:

- quality of the wire conductor or thickness of the wire
- height that the antenna is mounted above the ground
- capacitive end effects
- closeness to radio conductors

One simple method of partly improving antenna performance is by controlling the length of the antenna wire while measuring the Standing Wave Ratio (SWR) value, which is the ratio of the maximum peak voltage anywhere on a line to the minimum value at any point along the same line. The use of a SWR analyser becomes invaluable in this process.



Fig. 20: Measuring the SWR of the antenna with an SWR Analyser

The initial length of antenna wire was cut to match the length produced by the equation in Figure: 8 and was then cut in half before each end was connected to a dipole centre. The antenna was placed on the ground before being connected to the SWR analyser. The SWR value for the antenna was measured, and it was discovered the antenna was suitable for capturing of signals at 18 MHz. The antenna was then shortened by a small identical amount on each side and the

SWR values were again measured until the antenna registered as being most suitable for operation at 20.1MHz as seen in Figure: 20.

The SWR equations shown in Figure: 21 $|\rho|$ is the reflection coefficient, P_f is the power in the forward wave entering the antenna while P_r is the power of the reflected wave in the antenna [ARRL, 2000].

$$SWR = \frac{1 + |\rho|}{1 - |\rho|} \quad (7)$$

$$|\rho| = \frac{SWR - 1}{SWR + 1} \quad (8)$$

$$|\rho| = \sqrt{\frac{P_r}{P_f}} \quad (9)$$

Fig. 21: Standing Wave Ratio Equations

$$|\rho| = \frac{1.2 - 1}{1.2 + 1} \quad (10)$$

$$|\rho| = 0.0909090909090909088 \quad (11)$$

$$0.0909090909090909088 = \sqrt{\frac{P_r}{100}} \quad (12)$$

$$0.0909090909090909088^2 = \frac{P_r}{100} \quad (13)$$

$$P_r = 0.826446280991735 \quad (14)$$

Fig. 22: Calculating P_r for the antenna

The value of SWR for the antenna should be ideally as close as possible to 1, this ensures all the signal energy collected by the antenna is transferred to the load successfully. However in practice this is never the case as all antennas have varying amounts of loss. The antenna once mounted had a measured SWR value of 1.2 which equates to a P_r value as shown in Figure: 22 of 0.826446280991735 which is a loss of 17.35% or about 1dB which is considered reasonable [ARRL, 2000].

Table: 4 gives a breakdown of the costs associated with the antenna build to date.

Item Name	Cost (€)
100m Copper Wire	20.00
100m Coax RG59	25.00
2 × Dipole Centers	20.00
8 × PL259	18.00
2 × SO239	4.00
SMA-SO239 Adapter	4.50
2 × SO239-SO239 Adapter	4.50
T Adapter	3.00
Patch Lead	5.00
2 × 4m 5cm Pipe	15.00
2 × 4m 4cm Pipe	15.00
Total	134.00

Table 4: Antenna Build Cost

Listening Site Suitability

A convenient location for the antenna deployment was selected and the antenna was deployed. It was necessary to determine if this site was in fact a suitable location to host a radio telescope antenna. In order measure the ambient noise at the site, a Tektronics SA2600 spectrum analyser was connected to the antenna and left for 2 hours at midday to gather data. The spectrum is likely to be full of radio interference from human and natural sources at this time. The results can be seen in Figure: 23. The ambient RF noise at the site was quite high, ranging between -80dB and -70dB on average.

Similar readings were taken using the HackRF while attached to the antenna as show in Figure: 24. The measurements recorded with the HackRF were performed using the *gqrx* SDR software while all amplification settings inside *gqrx* were set at 0, in order to get an accurate reading on the raw signal being collected by the antenna. The values were on average -20dB quieter than those recorded by the SA2600 spectrum analyser, this points towards an inher-

ent deafness flaw in the HackRF design itself, and possibly the requirement of a preamplifier (preamp) in order to boost the gain on the input signal.



Fig. 23: Site Survey performed using a Tektronics SA2600 Spectrum Analyser



Fig. 24: Site Survey performed using HackRF One

In May 2013, the Solar Physics Group from Trinity College Dublin, performed a site survey at Birr Castle to ascertain the suitability of the site for an installation of the Low Frequency Array (LOFAR) telescope [McCauley, 2013]. The study was focused on a wide range of frequencies between 10 and 400MHz, a portion of which can be seen in Figure: 25. The report concluded that the location at Birr Castle was a relatively quiet location as compared against a site at Bornim, Germany and another at Bleien, Switzerland.

From the results of this survey, it can be seen that the average background noise levels were about -55dB near 20MHz which is similar to those recorded 80km away at the Kilkenny site. Table: 5 details the results collected at the

Kilkenny site and also includes those performed by the Solar Physics Group at Birr Castle for comparison. Results indicate the Kilkenny site appears to be a reasonable location at which to host the radio telescope.

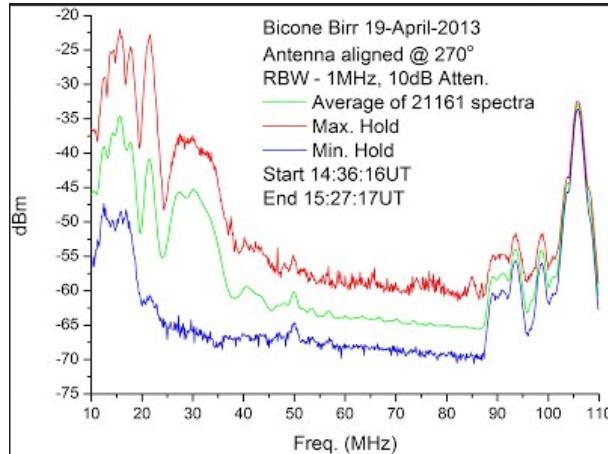


Fig. 25: LOFAR Site Suitability Survey performed at Birr Castle [McCauley, 2013]

Site	Equipment	Average Noise at 20.1MHz
Kilkenny	Tektronix SA2600 20.1MHz Dipole	-65dB
Kilkenny	HackRF One 20.1MHz Dipole	-65dB
Birr	LOFAR Array 10-100MHz Schwarzbeck Bicone	-55dB

Table 5: Average Noise Measurements at Kilkenny Listening Site compared to Birr Listening Site [McCauley, 2013]

Verification of the Radio Receiver

An opportunity to validate the HackRF transceiver arose with the announcement of a public Medium Wave (MW) experiment by Nichols [2015a] of the Radio Society of Great Britain (RSGB). The RSGB Propagation Studies Committee intended to perform an experiment during the partial solar eclipse on 20th March 2015. As can be seen in Figure: 49, the UK and Ireland would witness more than 90% totality during the eclipse. This was an opportunity to perform simple experiments to demonstrate the Sun's effect on Earth's ionosphere, and how this ionisation affects the propagation of radio signals.

According to Nichols [2015b], MW radio stations which are located more than 500 km away are unlikely to be propagated during daylight hours due to their signals being absorbed by the D layer as shown in Figure: 26. However this D layer does not exist during night time hours, as a result these radio signals are free to reflect from the E and F layers. Nichols [2015a] states that this is also true during a solar eclipse with a high totality percentage as would occur on the 20th March 2015.

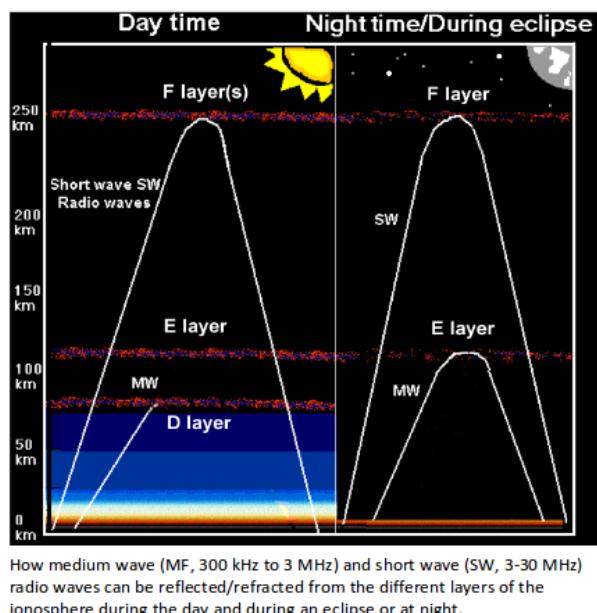


Fig. 26: Comparing propagation of MW radio signals during day and during night/solar eclipse [Nichols, 2015a]

Nichols [2015a] provided some a number of radio stations which transmit signals in the MW range at various locations in the United Kingdom, Europe and also a number in Iceland. The experiment called for choosing a station which is transmitting in the MW range and is also capable of being heard at the listeners location during the night time. This station should not be capable of being heard during the day time, the experiment called for a station which is further away than 500 km. The chosen radio station should then be tuned in to coincide with the beginning of the solar eclipse. The station receive power should be plotted against time. It was expected that the signal strength of the radio station should rise to a maximum which coincided with the solar eclipse maximum.

It was immediately apparent that the RadioJove antenna design would not be suitable for this experiment, and a suitable alternative was examined. During the initial research into potential designs a paper by Litva and Rook [1976] mentioned the use of a Beverage antenna for use in highly directional receivers such as early transatlantic radio communication experiments. Due to the simple nature of the Beverage antenna design, this was considered for use in this experiment.

Methodology

The following methodology was followed during this experiment:

1. The dipole antenna built for use in the main part of the dissertation was not best suited for picking up signals in the MW range so an alternative design was chosen.
2. A Beverage antenna was picked as it was easily aimed and had a simple design. The antenna very simply consisted of a copper wire 80 m long attached to a long wire antenna terminator which had an SO239 connector [Litva and Rook, 1976].
3. The listening site was longest in the East-West axis, and having the receiver positioned at the far west end of the site was most convenient. For this reason a radio station situation in an Easterly direction was picked.
4. The antenna was aimed in the general direction of the radio station chosen, and simply placed on the ground.
5. HackRF was tuned to the chosen radio station.

6. 2 MHz bandwidth of the spectrum was recorded during the entire length of the solar eclipse, the signal strength of the radio station can be plotted against time at a later date using this recorded data.
7. The SDR receiver should have the automatic gain control switch off in order to get accurate power results.

The chosen radio station was *Radio China International* (1.440 MHz) which was a powerful signal during the night time despite it being transmitted from Marnach in Luxembourg which is 950 km away from the listening site. Once the solar eclipse began, the HackRF SDR receiver was tuned to 1.440 MHz with a bandwidth of 2MHz and IQ samples were recorded. As the eclipse neared 20% totality, the radio station began to emerge from the background noise at -55 dB. The signal rose steadily until the eclipse maximum point which was -28 dB, and then averaged at a level of -45dB as the totality value dropped. At 20% of eclipse totality, the station faded once more before disappearing into the background noise level at -55dB. The results were plotted in Figure: 27 and were submitted to Nichols [2015a] for inclusion in the overall experiment.

A number of issues with the HackRF SDR transceiver were discovered during this experiment. The technical specifications show the operating range of the HackRF to be 10 MHz - 6 GHz [Ossmann, 2015d]. In practice it is possible to tune to frequencies below 4 MHz, even as low as 500kHz with a few caveats.

Reliable tuning proved a problem below 10MHz, when the HackRF was power cycled, tuning to the same frequency in software would see the station vanish. It took trial and error to find it once more by scanning lower and higher until the station was reacquired. This is due to the internal signal oscillator being affected by many external variables such as temperature. To be certain of the exact frequency being monitored, an external clock source must be used to provide a signal reference to the system.

The HackRF also appears to be quite deaf without maxing the IF and RF amplifiers. While this could be solved by adding a preamp at the antenna, such a step would also increase the level of background noise which might defeat the purpose.

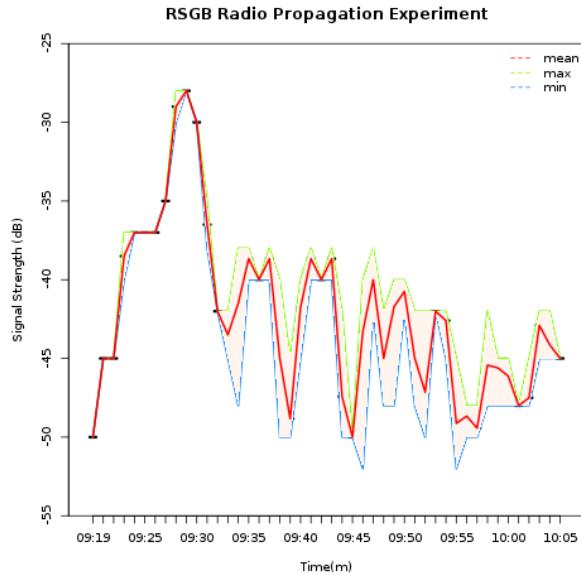


Fig. 27: Plot of signal power vs time during 20th March 2015 solar eclipse

Capturing DAM Emissions

Using the Radio Jupiter Pro software, a prediction was made that 23rd April 2015 would offer a very good opportunity to hear DAM emissions. As can be seen in Figure: 28, at the point of the highest chance of capturing DAM emissions, Jupiter was right on the very edge of the antenna beam focus. Due to the limitations in the design of the antenna, this configuration offered the best chance to receive an emission from Jupiter.

Despite this, a number of emissions were recorded as can be seen in the raw image Figure: 29. The emissions appear as faint horizontal lines in this waterfall diagram. This diagram is limited by the small bandwidth of 2 MHz and short period of 2 minutes. Ideally in order to spot these interesting emissions a large bandwidth of 20 MHz for instance should be used, over a long period of observation such as 1-24 hours. An example of this is section of a spectrogram seen in Figure: 30, which shows averaged signal power data for frequencies 17-55 MHz and 70-1000 MHz for a 24 hour period.

There are many open source software tools such as `rtl_power`, `simple_ra` and `heatmap` which could capture averaged raw signal strength data from a RTL based receiver in order to build spectrogram diagrams. Unfortunately most of

this software is not compatible with the HackRF. However the basic functionality from these tools can be replicated using basic processing blocks within GNURadio and the spectrograms can be built using tools like Octave or Matlab.

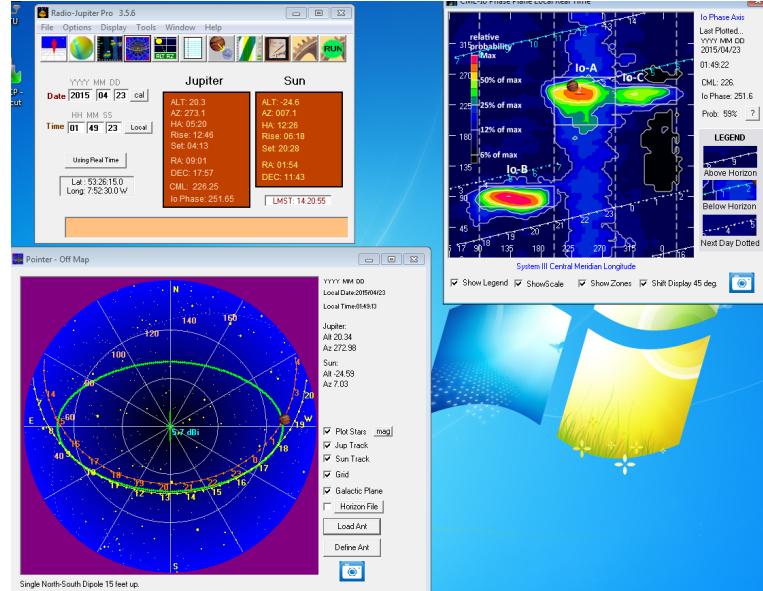


Fig. 28: Jupiter DAM IO-A Storm Prediction

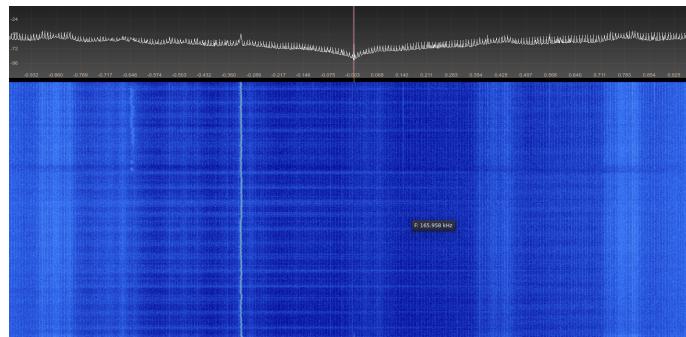


Fig. 29: DAM emissions picked up during an IO-A storm

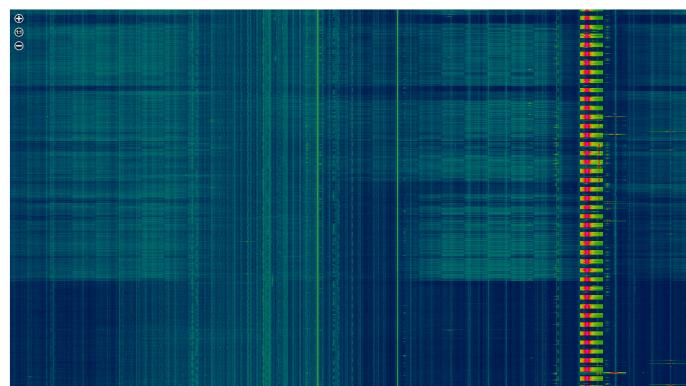


Fig. 30: Radio Spectrogram Diagram [Superkuh, 2007]

Research Question 1

What current Internet of Things (IOT) technologies would best suit the development of a software defined radio signal listening station and how cheaply can it be created?

- How expensive would a renewable power supply solution such as solar be?
- Can a low powered computing platform such as the Raspberry Pi B+ suitably host the system, or will a more powerful platform be required?
- How feasible are wireless technologies such as Zigbee, WIFI or Bluetooth to stream collected data to a central facility?

Determine SDRT Power Requirements

A major component of designing a self sufficient system is choosing the power supply. One solution is the use of renewable sources such as wind or solar to create energy and the use of batteries to act as a buffer and storage of this energy. When a suitably large battery array is deployed, power to the system can be provided during periods when there is no sunshine or wind available.

In order to design a self sufficient listening site it was necessary to firstly determine the power requirements of the hardware essential to the functioning of the system. If there are high computational requirements the system will generally also have high power needs, which can have a drastic effect on the cost of implementing a sustainable power supply using renewables. It is prudent therefore to implement the system as cheaply as possible, and using low powered devices.

The SDR RTL and HackRF receivers require a minimum of a USB 2.0 interface and computational requirements for processing data which immediately reduced the potential platforms to devices such as the Raspberry Pi B+, Raspberry Pi B2 and the Intel Atom or better.

Bryce [2011]

Methodology

The following platforms were chosen for testing purposes as they initially appear to meet or exceed the minimum requirements of operating *rtl_power* while also building spectrograms from the recorded data using *heatmap.py*:

1. Raspberry Pi B+ 800MHz, 512MB RAM
2. Raspberry Pi B 2. 900MHz, 1GB RAM
3. EEEPC Intel Atom N550 Netbook. 1.55GHz, 2GB RAM

Each platform would be tested on the following metrics:

Idle

The power usage while the system is at rest

heatmap.py

The power usage while the system is building spectrogram with heatmap.py

rtl_power

The power usage while the system is collecting data with rtl_power

Both heatmap.py and rtl_power

The power usage while the system is building spectrogram with heatmap.py and collecting data with rtl_power

100% CPU Load

The power usage while the system is at 100% load

The power usage of each system was measured using a power meter and the results were plotted as shown in Figures: 31, 32, and 33.

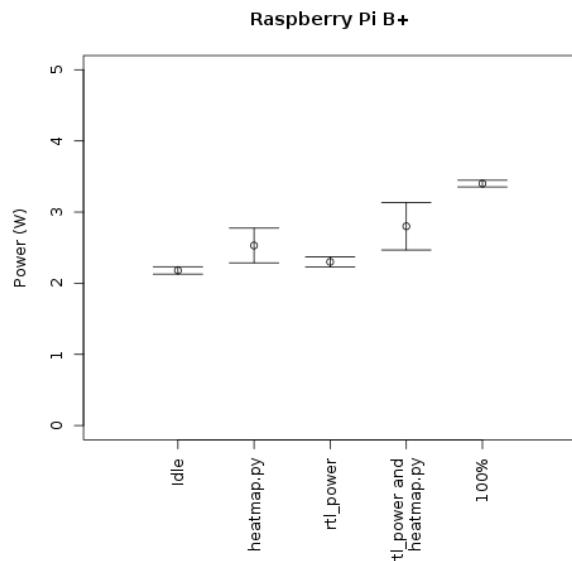


Fig. 31: SDRT Power Usage Raspberry Pi B+

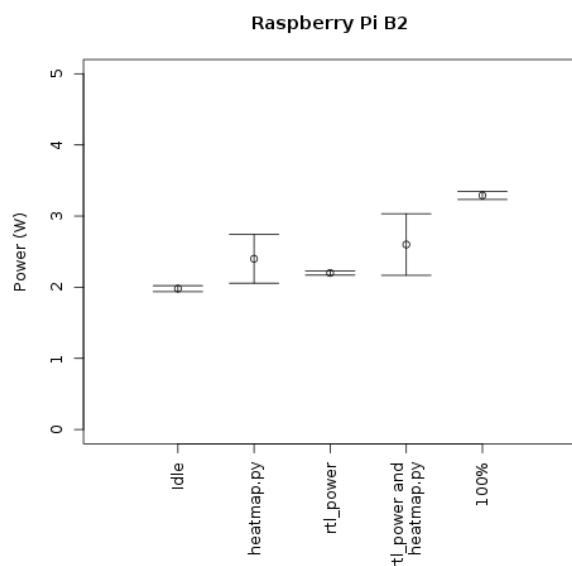


Fig. 32: SDRT Power Usage Raspberry Pi B2

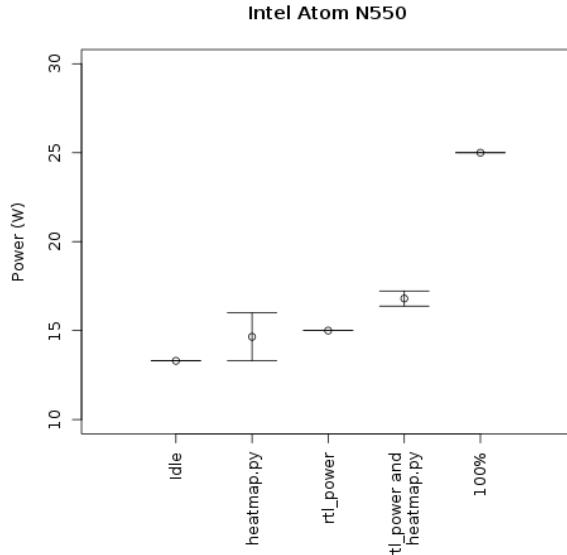


Fig. 33: SDRT Power Usage Intel Atom N550

Using solar data collected by the national weather service, MET Eireann [MET, 2015] and [ECA&D, 2015], a series of plots were produced to aid in the calculation of the battery array capacity required to power the system. As can be seen in the Figure: 37, the mean number of hours of sunlight for the year was calculated to be $\mu = 4.446$, with a minimum of 1 hour and a maximum of 9.156 hours.

$$\text{Power(Watts)} = \text{Current(Ampere)} \times \text{Voltage(V)} \quad (15)$$

$$A = \frac{W}{V} \quad (16)$$

$$A = \frac{25}{12} \quad (17)$$

$$\text{Ampere} = 2.08333333A \quad (18)$$

Fig. 34: Calculating Amps using Watts and Voltage

Solar Panels generally operate at either 12 or 24V, and when coupled with a battery charge controller, can be used to safely recharge batteries using generated power. The load power value measured previously for the Intel N550 Netbook operating at 100% was 25W. Taking the load and the assumed voltage value of 12V, the load current was generated using the power equation shown in Equation: 34. The current required to generate 25W at 12V is 2.0833333A.

An initial estimate of 48 hours of power capacity was made to ensure the system can be powered for an extended period during periods of low solar activity. A plot of load vs battery capacity was generated as shown in Figure: 36. This plot shows the capacity value required to satisfy a 25W load at 12V for a 48 hour period is about 100Ah.

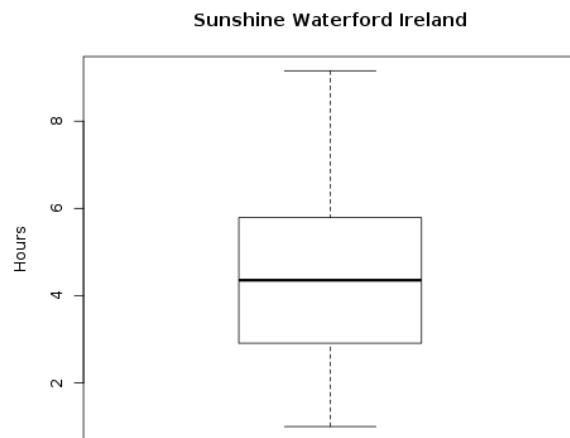


Fig. 35: Average Daily Sunshine in Waterford

Figure: 37 shows the averaged daily hours of sunshine over a 15 year period. This plot demonstrates that for half the year the sun shines for between 1 and 4 hours each day. This is a major limiting factor in creating a self sufficient solar powered system, as batteries may only be charged during periods of sunshine. During the times of the year with daily sunshine hours below the mean, a larger battery capacity is required in order to continue to power the system during

days with little sunshine. Correspondingly, as the sunshine hours fall, a larger solar panel is required to ensure as much light as possible is converted to energy storage.

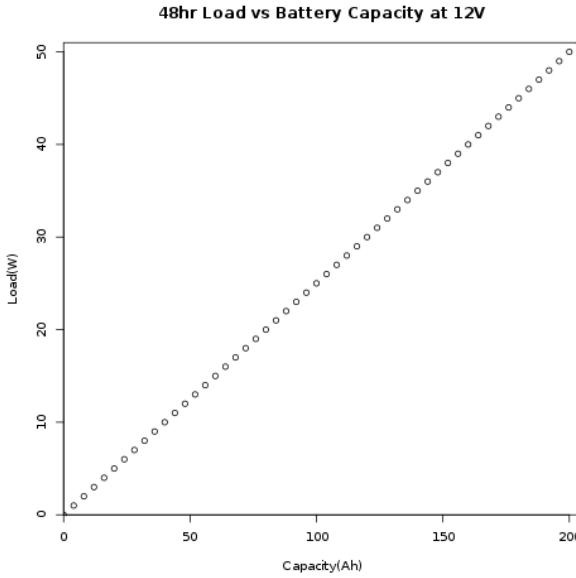


Fig. 36: Load vs Battery Capacity Plot

Next a plot of solar panel power output vs time taken to charge a 100Ah battery was generated as Figure: 38. This plot shows that in order to reliably charge a battery of size 100Ah in a single day which has the mean hours of sunshine the system will require a solar panel array of at least 350W.

The measurements were repeated for the Raspberry Pi systems. The power demands were rounded up to 5W which produced a battery capacity of 20Ah to power for 48 hours. The results were then tabulated as shown in Table: 6. The components required to implement a solar powered system such as the panels, batteries, charge controller and AC/DC inverters were priced and the results were included in this table also. Price wise, having a Raspberry Pi system in the listening station would make for a cost effective solution. As both the B+ and the B2 versions operate with the very low power usage of 5W, the solar renewable solution would be cheapest also at an estimated €150 for the solar power components.

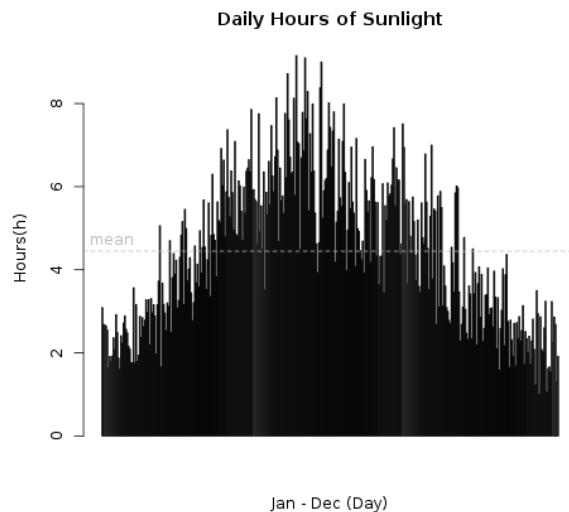


Fig. 37: Average Hours of Sunshine per Day in Waterford

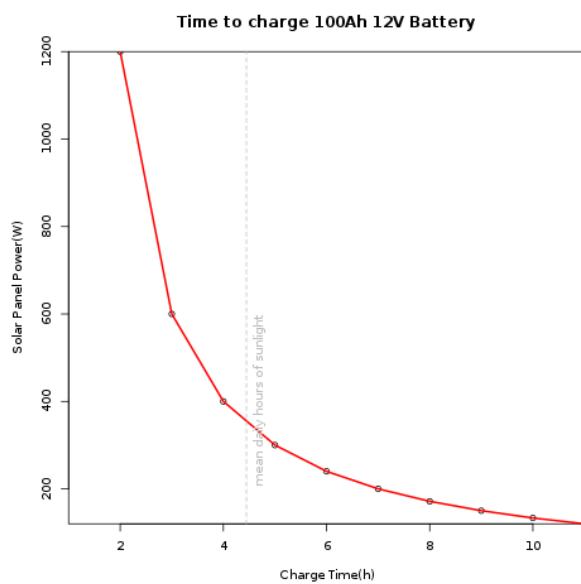


Fig. 38: Size of Solar Panel Required to Charge 100Ah Battery in Hours

Equipment	Solar Panel Power(W)	Battery Capacity(Ah)	Estimated Cost(€)
Raspberry Pi B+	75	20	150
Raspberry Pi B2	75	20	150
Intel Atom N550	350	100	500

Table 6: System Power Requirements

Determine SDRT Computational Power Requirements

The aim of this experiment was to determine the minimal hardware which was computationally capable of collecting data, while performing any data analytical or processing tasks required by the system.

Methodology

The following platforms were chosen for testing purposes as they initially appear to meet or exceed the minimum requirements of operating *rtl_power* while also building spectrograms from the recorded data using *heatmap.py*:

1. Raspberry Pi B+.
2. Raspberry Pi B 2.
3. EEEPC Intel Atom N550 Netbook.
4. Intel i5-3340M 2.7GHz

Each platform would be measured under a single metric, namely the length of time it takes to produce a spectrogram from data collected over a 24 hour period. Table: 7 shows the results gathered on each platform. As the time taken to process the spectrogram was longer than 24 hours for both Raspberry Pi models tested, they could not be considered capable of hosting the SDRT system and so this particular test was aborted when the time exceeded 24 hours for both Pi models.

The Intel Atom powered Netbook performed considerably better with a result of 4.12 hours. This would allow the collection of data, while also the sustainable processing of the previous days data without becoming overloaded.

Equipment	Generate Spectrogram(h)
Raspberry Pi B+ 800MHz 1 Core	24Hrs+
Raspberry Pi B2 900MHz 4 Core	24Hrs+
Intel Atom N550 1.55GHz 2 Core	4.12hours
Intel i5-3340M 2.7GHz 2 Core	0.25hours

Table 7: System CPU Power Requirements

Determine SDRT Network Connectivity Requirements

The aim of this experiment was to determine the optimal networking technology to transfer collected and or processed data from the SDRT to another system.

Methodology

The following technologies were investigated during this experiment:

1. Zigbee (XBEE Pro 2B 802.15.4)
2. WIFI (802.11N)
3. Fast Ethernet (802.3u)
4. Gigabit Ethernet (802.3-2008)

Each technology was measured using a single metric, transfer speed to copy 2GB of collected data between two nodes in the network. Table: 8 shows the results obtained.

Technology	Bitrate	Est(s)	Result(s)
Zigbee	250kbit/s	76,000	24Hrs+
WIFI N	144Mbit/s	132	737
Fast Ethernet	100Mbit/s	181	904
Gigabit Ethernet	1Gbit/s	30	45

Table 8: Networking technologies transferring 2GB data

There existed a large discrepancy between the estimated values and the actual measurements obtained for several technologies measured. A number of

factors could account for this, such as the estimated values being generated using theoretical maximum bandwidth for the technology, which are unrealistic for real world applications.

One example is the underlying hardware which handles both the USB and the Fast Ethernet network interfaces on the Raspberry Pi systems has a limit of 2Mbit/s, and could therefore never reach the theoretical maximum of 11Mbit/s suggested by Fast Ethernet. Using the Zigbee XBEE 802.15.4 modules, the theoretical estimate of 250kbit/s suggested a time of 21 hours to transfer 2GB of data, however after 24 hours the transfer was still ongoing. As the ability to the collected data must be faster than the time taken to collect it, the Zigbee test was cancelled after 24 hours. The results show WIFI, and Ethernet technologies being preferable in this instance.

Research Question 2

What processes or algorithms need to be developed to filter or flag known instances of human interference from radio signal observations?

- Flag transmission signals identified by amateur radio enthusiasts from a local DXSpider server in recorded data.
- Flag instances of natural radio interference such as lightning from the Blitzortung server in recorded data.

A DX server is a system which Amateur (Ham) Radio operators may use in order to inform one another in real time, about instances of radio signal transmissions from DX stations and other amateur radio stations. These DX Servers are often connected together to form a cluster. This allows many users from all over the world to communicate and share this information by connecting to the cluster. It is a rich source of identified human transmissions which often have the time, location and signal frequencies which were recorded. The data being shared is also often rigidly structured which greatly simplifies the task of implementing a software parsing solution [Koopman, 2007].

A software client was developed in order to connect to a local server within the DX Cluster. It was designed to gather all data from the cluster while filtering transmissions outside the frequency range currently being monitored by the telescope. All signals which have been identified as being within the range being monitored, would be flagged and all related meta data for each transmission would be stored for processing at a later stage.

The Blitzortung project operates an online service which displays global information about instances of lightning strikes in real time. Wanke et al. [2014] provide the designs for a listening station which amateur enthusiasts can then build and operate. Each site can then share the information back to a central server in real time [Wanke et al., 2014]. A software client could be developed which might connect to this Blitzortung server and retrieve occurrences of lightning strikes within a certain range such as 500 km centred on the listening site.

It is envisioned the time and frequency values associated with this data related to human signals, along with the data gathered regarding lightning strikes could be used to generate an overlay for a waterfall spectrogram diagram of the previous 24 hours worth of data collected by the telescope.

Experiment 1

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Experiment 2

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Experiment 3

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Research Question 3

What SDR processes and algorithms need to be developed to identify instances of the three main DAM emission types detailed in Table. 1?

- Can existing signal demodulation techniques be employed to find DAM emissions?
- What is required from an algorithm which might identify DAM emissions within an IQ signal?

There is a large amount of complexity involved in digital signal processing and the various techniques used to manipulate collected signals. This became apparent early during the literature review. An automated system which can automatically identify instances of Jovian DAM emissions would require some form of machine learning and this increases the complexity further.

Many websites dedicated to education and distance learning offer free 3rd level courses in topics related to those covered during this thesis. Coursera in particular offers a class on machine learning and a range of others in the fields of mathematics and even digital signal processing. It might prove beneficial to partake in one or more of these classes in the near future. A fully automated signal identification would seem to be far too complex to tackle during the scope of this dissertation, however some rudimentary machine learning techniques might yet prove useful in other areas of the system where decision making is required.

Ossmann [2015a] the developer of the HackRF SDR, offers a free course on the fundamentals of software defined radio, related mathematics such as complex number theory, trigonometry and digital signal processing. This has proved most useful in understanding how analogue signals are stored digitally, while also providing the necessary training in the basic principles of how frequency modulation and amplitude modulation for example is performed by a regular radio receiver.

As a large amount of work on amplitude modulation has already been completed for instance, it might be possible to use techniques for spotting AM signals in order to identify DAM emissions. More work is required to understand how this could be achieved, and might ultimately prove beyond the scope of what is achievable before the end of the dissertation.

Experiment 1

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Experiment 2

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Experiment 3

Methodology

The following methodology was followed during this experiment:

1. Aasdfasdf asdfasdfasdfasdf adsfasdf asdf s asdf asdfadsfasdf.

Software 1

Chapter for each software we produced how we validated it and proved it worked
tools and methodologies we used

Conclusions

In this interim report, the initial proposal has been reviewed, and current progress has been documented. At several points corrective modifications have been made to the scope. This ensures it remains possible to make progress within the time available.

In the initial proposal I conjectured at the potential pitfalls which would befall the project and hinder progress. I was not prepared for the inane problems which have so far caused the most trouble. Examples include the laptop USB ports being too slow to operate the HackRF at higher bandwidths. This forces me to sample at much narrower bandwidths which makes it very difficult to spot DAM emissions as they often are bigger than my window of the radio spectrum. Another issue I have encountered so far is the antenna mounting poles are too large to fit in the car. To address this I had to reduce size by 25%. This has reduced the potential configurations of the telescope to 10 ft, and 15 ft. This results in wasted observation opportunities as Jupiter can be at a point in the sky I cannot aim the telescope at.

Nor was I prepared for the complexity of digital signal processing and machine learning. I have taken mathematics refresher courses in order to understand the fundamentals of working with these signals. Several times, I made presumptions that I understood a concept only to discover it was far more complex which then added almost unbounded risk to the project. Due to issues encountered such as these, the scope has been changed drastically. I no longer expect to produce an automated identification algorithm to spot DAM emissions in recorded data. I expect an implementation might instead be completed in future work.

I am delighted to say I have lost no enthusiasm for the topic, after my experiences I am in fact encouraged to continue working on the topic further. I am considering applying to a doctoral program and hope pursue a research career in Astronomy.

Contribution to Research Knowledge Anticipated

The contributions to research knowledge which are anticipated are listed as follows:

1. A low cost radio telescope listening platform design which Amateur astronomers can use to collaborate on radio astronomical research.
2. A rudimentary SDR solution to filter or flag human and natural radio interference from observations.
3. An SDR solution to produce waterfall spectrogram diagrams with interference sources flagged in an overlay.
4. Jovian and or Solar radio emissions such as those listed in Table: 1 may become apparent when a large bandwidth and long period spectrogram is created.

Future Work

A dissertation is a work of scholarly investigation that is grounded in the research literature and differs from a report or a book. It is judged on a prescribed set of academic criteria. Although the likely outcomes are tentative at the start of the program, it is useful to incorporate them into the research proposal to help focus the work program.

Here is something which can be continued in future work

Appendix

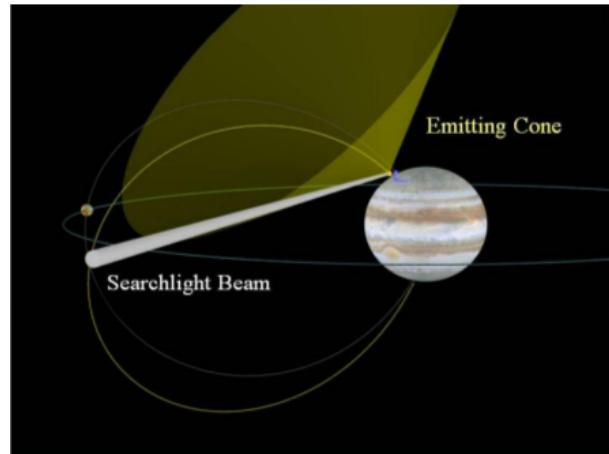


Fig. 39: A searchlight beam model of Jupiter's decametric radio emissions [Imai et al., 2008]

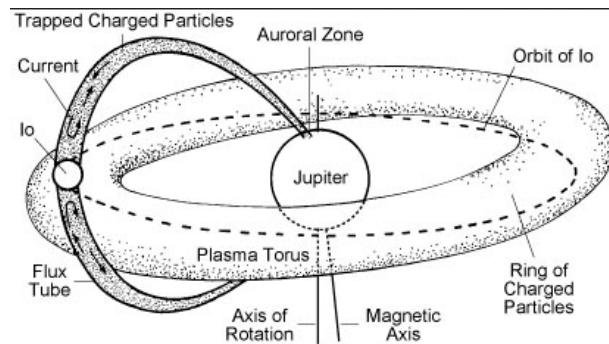


Fig. 40: Io Flux Tube and the Plasma Torus [Lang, 2010]

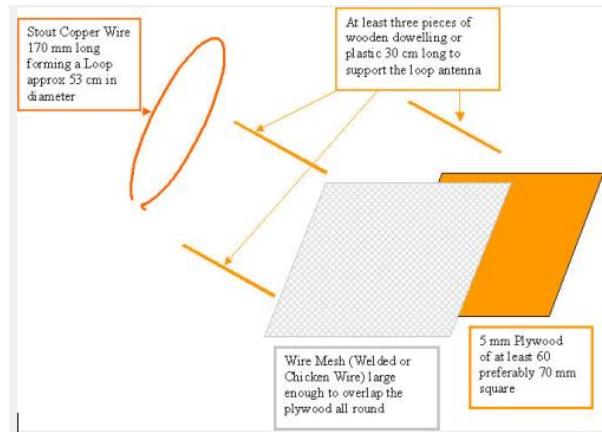


Fig. 41: 21 MHz Shortwave Loop Antenna Radio Telescope Design [Greef, 2012]

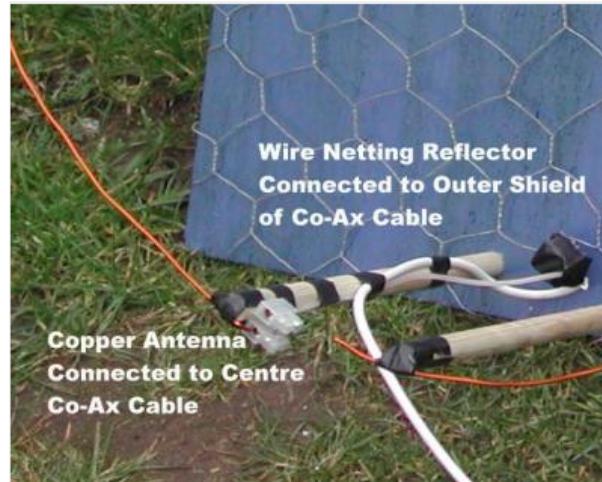


Fig. 42: 21 MHz Shortwave Loop Antenna Radio Telescope Design [Greef, 2012]

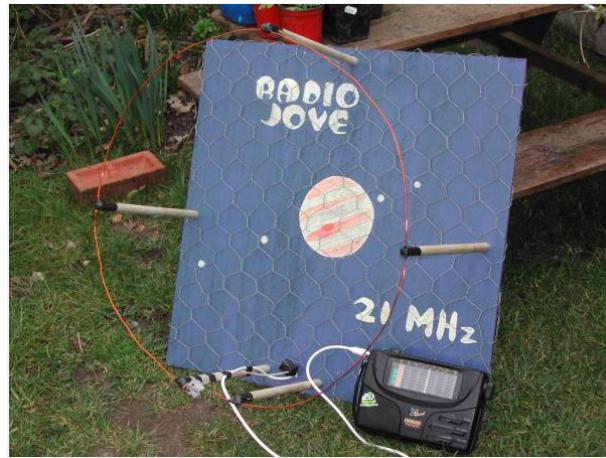


Fig. 43: 21 MHz Shortwave Loop Antenna Radio Telescope Design [Greef, 2012]



Fig. 44: Dual Dipole Antenna Deployment

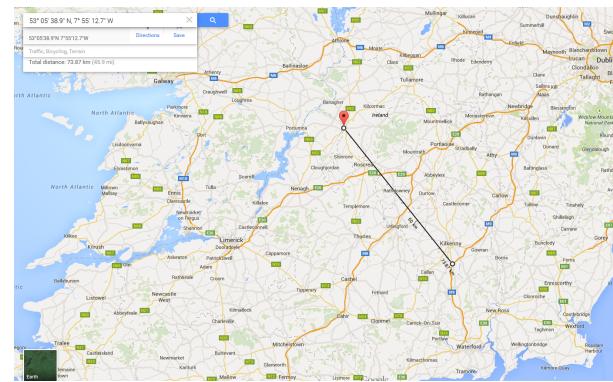


Fig. 45: Distance between the Birr and Kilkenny site survey

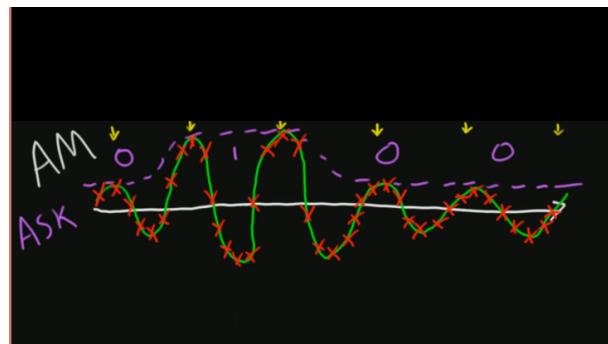


Fig. 46: Sampling an Analogue Amplitude Modulation Signal [Ossmann, 2015c]

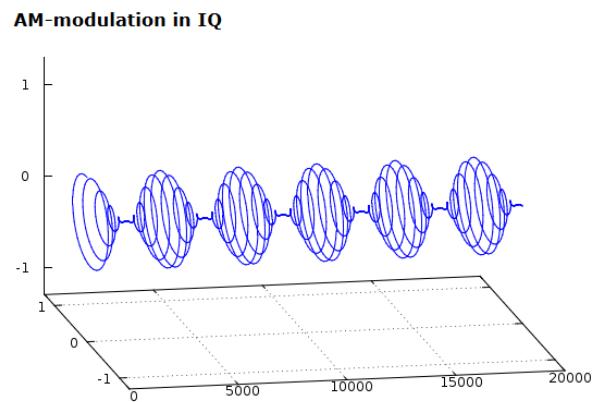


Fig. 47: I/Q Sample Polar Plot of an AM Modulated Signal [Kuisma, 2014]

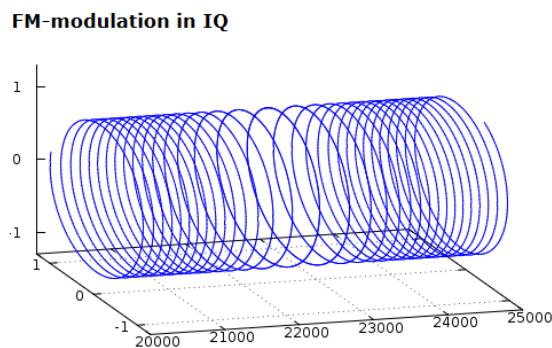


Fig. 48: I/Q Sample Polar Plot of an FM Modulated Signal [Kuisma, 2014]

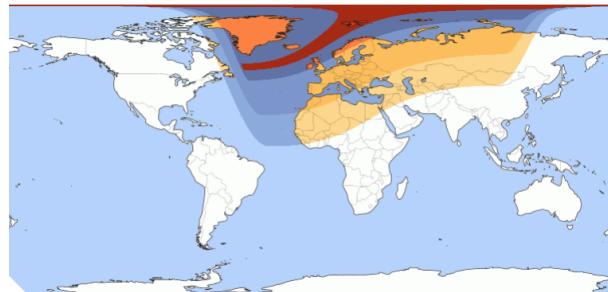


Fig. 49: Partial Solar Eclipse 20th March 2015

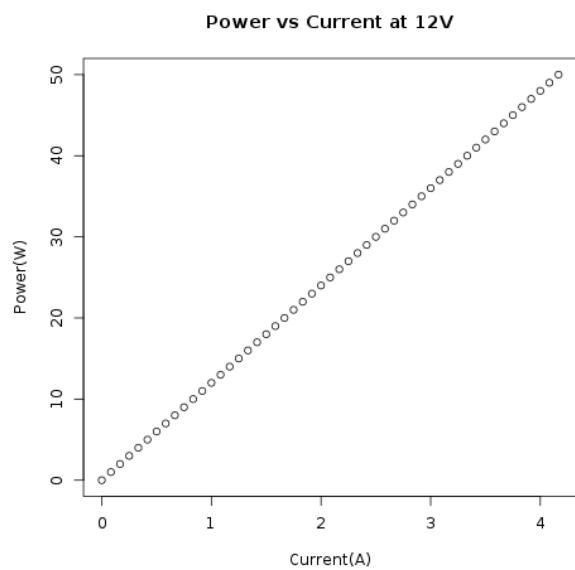


Fig. 50: Power vs Current at 12V Plot

Bibliography

- ARRL. *ARRL Operating Manual 7th Edition*, 2000. URL <http://www.arrl.org/shop/ARRL-Operating-Manual-10th-Edition/>.
- Thomas Ashcraft. Jupiter 27 september 2013 n event, September 2013. URL http://www.heliotown.com/Jupiter_Sep272013_N_Event.html.
- John W. Belcher. The jupiter-io connection: An alfvén engine in space. *Science*, 238(Oct 9):170–176, 1987.
- S. K. Bose, S. Sarkar, and A. B. Bhattacharyya. Jovian decametric radio emission: An overview of the planetary radio astronomical observations. *Indian Journal of Radio and Space Physics*, 37:77–108, 2008.
- Michael Bryce. *Emergency Power for Radio Communications*, chapter 3, page 12. ARRL - The national association for amateur radio, 2011. ISBN 978-0-87259-615-3. URL <http://www.arrl.org>.
- F.B. Burke and L. K. Franklin. Obervations of variable radio source associated with the planet jupiter. *Journal of Geophysical Research*, 60(213), 1955.
- Comreg. *Overview Comreg - Radio Spectrum*. Commission for Communications Regulation, 2014. URL http://www.comreg.ie/radio_spectrum/overview.496.html.
- ECA&D. European climate assessment and dataset, 2015. URL <http://eca.knmi.nl/dailydata/customquery.php>.
- J. M. Freidt. Gnuradio as a digital signal processing environment: Application to acoustic wireless sensor measurement and time and frequency analysis of periodic signals. *Joint UFFC EFTF and PFM Symposium*, pages 278–281, 2013.
- Gnuradio. Tutorials - gnuradio, 2014. URL <http://gnuradio.org/redmine/projects/gnuradio/wiki/Tutorials>.
- B. Greef. A simple decametric radio telescope for jupiter, 2012. URL http://www.ukaranet.org.uk/uk_amateurs/bobgreef/.
- K. Imai, L. Garcia, F. Reyes, M. Imai, J. R. Thieman, and M. Ikuta. A Searchlight Beam Model of Jupiter’s Decametric Radio Emissions. *AGU Fall Meeting Abstracts*, page B1673, December 2008.
- M. G. Kivelson, K. K. Khurana, R. J. Walker, C. T. Russell, J. A. Linker, D. J. Southwood, and C. Polanskey. A magnetic signature at io: Inital report from the galileo magnetometer. *Science*, 273, July 1996.

- Dirk Koopman. Dxspider, 2007. URL <http://www.dxcluster.org/main/>.
- Mikael Q Kuisma. *I/Q Data for Dummies*, 2014. URL <http://whiteboard.ping.se/SDR/IQ>. Complex Numbers in DSP.
- K. R. Lang. Io flux tube and plasma torus, 2010. URL http://ase.tufts.edu/cosmos/view_picture.asp?id=1174.
- J Litva and B.J. Rook. *Beverage Antennas for HF Communications, Direction Finding and Over-the-horizon Radars*, chapter 1, pages 1–3. Communications Research Centre, 1976. URL <http://www.nrccdxas.org/articles/BevAnt0876.pdf>.
- Joe McCauley. Craf the esf expert committee on radio astronomy frequencies, 2013. URL http://www.craf.eu/newsletters/newsletter27_0713.pdf.
- MET. Download historical climate data for our synoptic stations, 2015. URL <http://www.met.ie/climate-request/>.
- NASA. *The Effects of Earth's Upper Atmosphere on Radio Signals*, 2012a. URL <http://radiojove.gsfc.nasa.gov/education/educ/radio/tran-rec/exerc/iono.htm>.
- NASA. *Antenna Kit and Manual Developed for NASA Radio JOVE Project by The Radio JOVE Project Team*, 2012b. URL http://radiojove.gsfc.nasa.gov/telescope/ant_manual.pdf.
- Steve Nichols. Partial solar eclipse propagation experiment, march 20th 2015, 2015a. URL <http://forums.thersgb.org/index.php?threads/partial-solar-eclipse-propagation-experiment-march-20th-2015.125/>.
- Steve Nichols. Partial solar eclipse propagation experiment, march 20th 2015, 2015b. URL <http://www.g0kya.blogspot.ie/2015/02/partial-solar-eclipse-propagation.html>.
- Michael Ossmann. *Software Defined Radio with HackRF*, 2015a. URL <http://greatscottgadgets.com/sdr/>.
- Michael Ossmann. *Software Defined Radio with HackRF*, 2015b. URL <http://greatscottgadgets.com/sdr/9>. Aliasing.
- Michael Ossmann. *Software Defined Radio with HackRF*, 2015c. URL <http://greatscottgadgets.com/sdr/7>. Complex Numbers in DSP.
- Michael Ossmann. *Software Defined Radio with HackRF*, 2015d. URL <http://greatscottgadgets.com/hackrf/>. HackRF One.
- Radio-Sky-Publishing. *Radio Jupiter Pro*, 2014. URL <http://www.radiosky.com/rjpro3ishere.html>.

- RSGB. *RSGB Radio Communication Handbook*, 2014. URL http://www.rsgbshop.org/acatalog/Online_Catalogue_Technical_6.html.
- Steven W. Smith. *Digital Signal Processing A Practical Guide for Engineers and Scientists*, chapter 1, pages 1–3. Newnes, 2003a. ISBN 0-75067444-X. What is DSP.
- Steven W. Smith. *Digital Signal Processing A Practical Guide for Engineers and Scientists*, chapter 3, pages 39–44. Newnes, 2003b. ISBN 0-75067444-X. Sampling Theorem.
- Superkuh. Radio spectrum, 2007. URL http://superkuh.com/radio/2014-08-26_17-55_70-1000_25k/.
- E. Wanke, R. Andersen, and T. Volgnandt. *A WorldWide Low Cost Community Based Time of Arrival Lightning Detection and Lightning Location Network*, 2014. URL http://www.blitzortung.org/Documents/T0A_Blitzortung_RED.pdf.
- M. Wilkinson and J. Kennewell. A project to detect radio emissions from jupiter. *Southern Sky*, 1994. URL <http://www.spaceacademy.net.au/spacelab/projects/jovrad/jovrad.htm4>.