

1 **The Sensitivity to Oscillation Parameters**
2 **from a Simultaneous Beam and**
3 **Atmospheric Neutrino Analysis that**
4 **combines the T2K and SK Experiments**



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Abstract

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Acknowledgements

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Introduction

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Neutrino Oscillation Physics

¹²⁹ When first proposed, neutrinos were expected to be approximately massless
¹³⁰ fermions that only interact through weak and gravitational forces. This meant
¹³¹ they were very difficult to detect as they can pass through significant amounts
¹³² of matter without interacting. Despite this, experimental neutrino physics has
¹³³ developed many different detection techniques and observed neutrinos from
¹³⁴ both natural and artificial sources. In direct tension with Standard Model physics,
¹³⁵ neutrinos have been determined to oscillate between different lepton flavours,
¹³⁶ requiring them to have mass.

¹³⁷ The observation techniques which led to the discovery of the neutrino are doc-
¹³⁸ umented in section 2.1. The theory underpinning neutrino oscillation is described
¹³⁹ in section 2.2 and includes the approximations which can be made to simplify
¹⁴⁰ the understanding of neutrino oscillation in the two-flavour approximation. Past,
¹⁴¹ current, and future neutrino experiments are detailed in section 2.3, including the
¹⁴² reactor, atmospheric, and long-baseline accelerator neutrino sources that have
¹⁴³ been used to successfully constrain oscillation parameters. Finally, the current
¹⁴⁴ state of oscillation parameter measurements are summarised in section 2.4.

2.1 Discovery of Neutrinos

At the start of the 20th century, the electrons emitted from the β -decay of the nucleus were found to have a continuous energy spectrum [1, 2]. This observation seemingly broke the energy conservation invoked within that period's nuclear models. In 1930, Pauli provided a solution to this problem in the form of a new particle, the neutrino (originally termed "neutron"). It was theorized to be an electrically neutral spin-1/2 fermion with a mass smaller than that of the electron [3]. This neutrino was emitted with the electron in β -decay to alleviate the apparent breaking of energy conservation. As a predecessor of today's weak interaction model, Fermi's theory of β -decay developed the understanding by coupling the four constituent particles: electron, proton, neutron, and neutrino, into a quantitative model [4].

Whilst Pauli was not convinced of the ability to detect neutrinos, the first observations of the particle were made in the mid-1950s when neutrinos from a reactor were observed via the inverse β -decay (IBD) process, $\bar{\nu}_e + p \rightarrow n + e^+$ [5, 6]. The detector consisted of two parts: a neutrino interaction medium and a liquid scintillator. The interaction medium was built from two water tanks, loaded with cadmium chloride to allow for increased efficiency in the detection of neutron capture. The positron emitted from IBD annihilates, $e^+ + e^- \rightarrow 2\gamma$, generating a prompt signal and the neutron is captured on the cadmium via $n + ^{108}Cd \rightarrow ^{109*}Cd \rightarrow ^{109}Cd + \gamma$, producing a delayed signal. An increase in the coincidence rate was observed when the reactor was operating which was interpreted as interactions from neutrinos generated in the reactor.

After the discovery of the ν_e , the question of how many flavours of neutrino exist was asked. In 1962, a measurement of the ν_μ was conducted at the Brookhaven National Laboratory [7]. A proton beam was directed at a beryllium target, generating pions which then decayed via $\pi^\pm \rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu)$, and the subsequent interactions of the ν_μ were observed. As the subsequent interaction of the neutrino generated muons rather than electrons, it was determined that

¹⁷⁴ the ν_μ was fundamentally different from ν_e . The final observation to be made
¹⁷⁵ was that of the ν_τ from the DONUT experiment [8]. Three neutrinos seem the
¹⁷⁶ obvious solution as it mirrors the known number of charged leptons (as they form
¹⁷⁷ weak isospin doublets) but there could be evidence of more. Several neutrino
¹⁷⁸ experiments have found anomalous results [9, 10] which could be attributed
¹⁷⁹ to “sterile” neutrinos. These hypothesised particles are not affected by gauge
¹⁸⁰ interactions in the Standard Model so their presence can only be inferred through
¹⁸¹ the observation of non-standard oscillation modes. However, cosmological
¹⁸² observations indicate the number of neutrino species $N_{eff} = 2.99 \pm 0.17$ [11], as
¹⁸³ measured from the cosmic microwave background power spectrum. LEP also
¹⁸⁴ measured the number of active neutrino flavours to be $N_\nu = 2.9840 \pm 0.0082$ [12]
¹⁸⁵ from measurements of the Z-decay width, but this does not strongly constrain
¹⁸⁶ the number of sterile neutrinos.

¹⁸⁷ 2.2 Theory of Neutrino Oscillation

¹⁸⁸ A neutrino generated with lepton flavour α can change into a different lepton
¹⁸⁹ flavour β after propagating some distance. This phenomenon is called neutrino
¹⁹⁰ oscillation and requires that neutrinos must have a non-zero mass. This behaviour
¹⁹¹ has been characterised by the Pontecorvo-Maki-Nakagawa-Sakata (PMNS) [13–
¹⁹² 15] mixing matrix which describes how the flavour and mass of neutrinos are
¹⁹³ associated. This is analogous to the Cabibbo-Kobayashi-Maskawa (CKM) [16]
¹⁹⁴ matrix measured in quark physics.

¹⁹⁵ 2.2.1 Three Flavour Oscillations

¹⁹⁶ The PMNS parameterisation defines three flavour eigenstates, ν_e , ν_μ and ν_τ
¹⁹⁷ (indexed ν_α), which are eigenstates of the weak interaction and three mass
¹⁹⁸ eigenstates, ν_1 , ν_2 and ν_3 (indexed ν_i). Each mass eigenstate is the superposition

¹⁹⁹ of all three flavour states,

$$|\nu_i\rangle = \sum_{\alpha} U_{\alpha i} |\nu_{\alpha}\rangle. \quad (2.1)$$

²⁰⁰ Where U is the 3×3 PMNS matrix which is unitary and connects the mass
²⁰¹ and flavour eigenstates.

²⁰² The weak interaction, when interacting via a W^{\pm} boson, couples to flavour
²⁰³ eigenstates so neutrinos interact with leptons of the same flavour. The prop-
²⁰⁴ agation of a neutrino flavour eigenstate, in a vacuum, can be re-written as a
²⁰⁵ plane-wave solution to the time-dependent Schrödinger equation,

$$|\nu_{\alpha}(t)\rangle = \sum_i U_{\alpha i}^{*} |\nu_i\rangle e^{-i\phi_i}. \quad (2.2)$$

²⁰⁶ The ϕ_i term can be expressed in terms of the energy, E_i , and magnitude of the
²⁰⁷ three momenta, p_i , of the neutrino, $\phi_i = E_i t - p_i x$ (t and x being time and position
²⁰⁸ coordinates). The probability of observing a neutrino of flavour eigenstate β from
²⁰⁹ one which originated as flavour α can be calculated as,

$$P(\nu_{\alpha} \rightarrow \nu_{\beta}) = |\langle \nu_{\beta} | \nu_{\alpha}(t) \rangle|^2 = \sum_{i,j} U_{\alpha i}^{*} U_{\beta i} U_{\alpha j} U_{\beta j}^{*} e^{-i(\phi_j - \phi_i)}. \quad (2.3)$$

²¹⁰ The term within the exponential can be represented as,

$$\phi_j - \phi_i = E_j t - E_i t - p_j x + p_i x. \quad (2.4)$$

²¹¹ For a relativistic particle, $E_i \gg m_i$, a Taylor series expansion means,

$$p_i = \sqrt{E_i^2 - m_i^2} \approx E_i - \frac{m_i^2}{2E_i}. \quad (2.5)$$

²¹² Making the approximations that neutrinos are relativistic, the mass eigenstates
²¹³ were created with the same energy and that $x = L$, where L is the distance
²¹⁴ travelled by the neutrino, Equation 2.4 then becomes

$$\phi_j - \phi_i = \frac{\Delta m_{ij}^2 L}{2E}, \quad (2.6)$$

where $\Delta m_{ij}^2 = m_i^2 - m_j^2$. This, combined with further use of unitarity relations results in Equation 2.3 becoming

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\beta) &= \delta_{\alpha\beta} - 4 \sum_{i>j} \Re \left(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin^2 \left(\frac{\Delta m_{ij}^2 L}{4E} \right) \\ &\quad + (-) 2 \sum_{i>j} \Im \left(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin \left(\frac{\Delta m_{ij}^2 L}{2E} \right). \end{aligned} \quad (2.7)$$

Where $\delta_{\alpha\beta}$ is the Kronecker delta function and the negative sign on the last term is included for the oscillation probability of antineutrinos. As an important point to note, the observation of oscillation probability requires a non-zero value of Δm_{ij}^2 , which in turn requires that neutrinos have differing masses.

Typically, the PMNS matrix is parameterised into three mixing angles, a charge parity (CP) violating phase δ_{CP} , and two Majorana phases $\alpha_{1,2}$,

$$U = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\text{Atmospheric, Accelerator}} \underbrace{\begin{pmatrix} c_{13} & 0 & s_{13} e^{-i\delta_{CP}} \\ 0 & 1 & 0 \\ -s_{13} e^{-i\delta_{CP}} & 0 & c_{13} \end{pmatrix}}_{\text{Reactor, Accelerator}} \times \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Reactor, Solar}} \underbrace{\begin{pmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Majorana}}. \quad (2.8)$$

Where $s_{ij} = \sin(\theta_{ij})$ and $c_{ij} = \cos(\theta_{ij})$. The oscillation parameters are often grouped: (1,2) as “solar”, (2,3) as “atmospheric” and (1,3) as “reactor”. Many neutrino experiments aim to measure the PMNS parameters from a wide array of origins, as is the purpose of this thesis.

The Majorana phase, $\alpha_{1,2}$, included within the fourth matrix in Equation 2.8 is only included for completeness. For an oscillation analysis experiment, any terms containing this phase disappear due to taking the expectation value of the PMNS matrix. Measurements of these phases can be performed by experiments searching for neutrino-less double β -decay [17].

232 A two-flavour approximation can be obtained when one assumes the third
233 mass eigenstate is degenerate with another. This results in the two-flavour
234 approximation being reasonable for understanding the features of the oscillation.
235 In this two-flavour case, the mixing matrix becomes,

$$U_{2\text{ Flav.}} = \begin{pmatrix} \cos(\theta) & \sin(\theta) \\ -\sin(\theta) & \cos(\theta) \end{pmatrix}. \quad (2.9)$$

236 This culminates in the oscillation probability,

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\alpha) &= 1 - \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right), \\ P(\nu_\alpha \rightarrow \nu_\beta) &= \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right). \end{aligned} \quad (2.10)$$

237 Where $\alpha \neq \beta$. For a fixed neutrino energy, the oscillation probability is
238 a sinusoidal function depending upon the distance over which the neutrino
239 propagates. The frequency and amplitude of oscillation are dependent upon
240 $\Delta m^2/4E$ and $\sin^2 2\theta$, respectively. The oscillation probabilities presented thus far
241 assume $c = 1$, where c is the speed of light in a vacuum. In more familiar units, the
242 maximum oscillation probability for a fixed value of θ is given at $L[\text{km}] / E[\text{GeV}] \sim$
243 $1.27/\Delta m^2$. It is this calculation that determines the best L/E value for a given
244 experiment to be designed around for measurements of a specific value of Δm^2 .

245 2.2.2 The MSW Effect

246 The theory of neutrino oscillation in a vacuum has been described in subsec-
247 tion 2.2.1. However, the beam neutrinos and atmospheric neutrinos originating
248 from below the horizon propagate through the matter in the Earth. The coherent
249 scattering of neutrinos from a material target modifies the Hamiltonian of the
250 system which results in a change in the oscillation probability. This modification
251 is termed the Mikheyev-Smirnov-Wolfenstein (MSW) effect [18–20]. This occurs
252 because charged current scattering ($\nu_e + e^- \rightarrow \nu_e + e^-$, propagated by a W boson)
253 only affects electron neutrinos whereas the neutral current scattering ($\nu_l + l^- \rightarrow$

²⁵⁴ $\nu_l + l^-$, propagated by a Z^0 boson) interacts through all neutrino flavours equally.
²⁵⁵ In the two-flavour approximation, the effective mixing parameter becomes

$$\sin^2(2\theta) \rightarrow \sin^2(2\theta_m) = \frac{\sin^2(2\theta)}{(A/\Delta m^2 - \cos(2\theta))^2 + \sin^2(2\theta)}, \quad (2.11)$$

²⁵⁶ where $A = 2\sqrt{2}G_F N_e E$, N_e is the electron density of the medium and G_F
²⁵⁷ is Fermi's constant. It is clear that there exists a value of $A = \Delta m^2 \cos(2\theta)$ for
²⁵⁸ $\Delta m^2 > 0$, which results in a divergent mixing parameter, colloquially called the
²⁵⁹ matter resonance. This resonance regenerates the electron neutrino component of
²⁶⁰ the neutrino flux [18–20]. The density at which the resonance occurs is given by

$$N_e = \frac{\Delta m^2 \cos(2\theta)}{2\sqrt{2}G_F E}. \quad (2.12)$$

²⁶¹ At densities lower than this critical value, the oscillation probability will
²⁶² be much closer to that of vacuum oscillation. For antineutrinos, $N_e \rightarrow -N_e$
²⁶³ [21]. The resonance occurring from the MSW effect depends on the sign of Δm^2 .
²⁶⁴ Therefore, any neutrino oscillation experiment which observes neutrinos and
²⁶⁵ antineutrinos which have propagated through matter can have some sensitivity
²⁶⁶ to the ordering of the neutrino mass eigenstates.

²⁶⁷ 2.3 Neutrino Oscillation Measurements

²⁶⁸ As evidence of beyond Standard Model physics, the 2015 Nobel Prize in Physics
²⁶⁹ was awarded to the Super-Kamiokande (SK) [22] and Sudbury Neutrino Ob-
²⁷⁰ servatory (SNO) [23] collaborations for the first definitive observation of solar
²⁷¹ and atmospheric neutrino oscillation [24]. Since then, the field has seen a wide
²⁷² array of oscillation measurements from a variety of neutrino sources. As seen
²⁷³ in subsection 2.2.1, the neutrino oscillation probability is dependent on the ratio
²⁷⁴ of the propagation baseline, L , to the neutrino energy, E . It is this ratio that
²⁷⁵ determines the type of neutrino oscillation a particular experiment is sensitive to.

²⁷⁶ As illustrated in Figure 2.1, there are many neutrino sources that span a
²⁷⁷ wide range of energies. The least energetic neutrinos are from reactor and

²⁷⁸ terrestrial sources at $O(1)$ MeV whereas the most energetic neutrinos originate
²⁷⁹ from atmospheric and galactic neutrinos of $> O(1)$ TeV.



Figure 2.1: The electro-weak cross-section for $\bar{\nu}_e + e^- \rightarrow \bar{\nu}_e + e^-$ scattering on free electrons from various natural and man-made neutrino sources, as a function of neutrino energy. Taken from [25]

²⁸⁰ 2.3.1 Solar Neutrinos

²⁸¹ Solar neutrinos are emitted from fusion reaction chains at the centre of the Sun.
²⁸² The solar neutrino flux, given as a function of neutrino energy for different
²⁸³ fusion and decay chains, is illustrated in Figure 2.2. Whilst proton-proton fusion
²⁸⁴ generates the largest flux of neutrinos, the neutrinos are low energy and are
²⁸⁵ difficult to reconstruct due to the IBD interaction threshold of 1.8MeV [26].
²⁸⁶ Consequently, most experiments focus on the neutrinos from the decay of 8B
²⁸⁷ (via $^8B \rightarrow ^8Be^* + e^+ + \nu_e$), which are higher energy.

²⁸⁸ The first measurements of solar neutrinos observed a significant reduction in
²⁸⁹ the event rate compared to predictions from the Standard Solar Model [28, 29]. A
²⁹⁰ proposed solution to this “solar neutrino problem” was $\nu_e \leftrightarrow \nu_\mu$ oscillations in a



Figure 2.2: The solar neutrino flux as a function of neutrino energy for various fusion reactions and decay chains as predicted by the Standard Solar Model. Taken from [27].

291 precursory version of the PMNS model [30]. The Kamiokande [31], Gallex [32]
 292 and Sage [33] experiments confirmed the ~ 0.5 factor deficit of solar neutrinos.

293 The conclusive solution to this problem was determined by the SNO col-
 294 laboration [34]. Using a deuterium water target to observe 8B neutrinos, the
 295 event rate of charged current (CC), neutral current (NC), and elastic scattering
 296 (ES) interactions (Given in Equation 2.13) was simultaneously measured. CC
 297 events can only occur for electron neutrinos, whereas the NC channel is agnostic
 298 to neutrino flavour, and the ES reaction has a small excess sensitivity for the
 299 detection of electron neutrino interactions. This meant that there were direct
 300 measurements of the ν_e and ν_x neutrino flux. It was concluded that the CC and
 301 ES interaction rates were consistent with the deficit previously observed. Most
 302 importantly, the NC reaction rate was only consistent with the others under the

303 hypothesis of flavour transformation.

$$\begin{aligned} \nu_e + d &\rightarrow p + p + e^- & (CC) \\ \nu_x + d &\rightarrow p + n + \nu_x & (NC) \\ \nu_x + e^- &\rightarrow \nu_x + e^- & (ES) \end{aligned} \quad (2.13)$$

304 Since the SNO measurement, many experiments have since measured the
 305 neutrino flux of different interaction chains within the sun [35–37]. The most
 306 recent measurement was that of CNO-cycle neutrinos which were recently
 307 observed with 5σ significance by the Borexino collaboration [35].

308 2.3.2 Accelerator Neutrinos

309 The concept of using an artificial “neutrino beam” was first realised in 1962 [38].
 310 Since then, many experiments have adopted the same fundamental concepts.
 311 Typically, a proton beam is aimed at a target producing charged mesons that
 312 decay to neutrinos. The mesons can be sign-selected by the use of magnetic
 313 focusing horns to generate a neutrino or antineutrino beam. Pions are the primary
 314 mesons that decay and depending on the orientation of the magnetic field, a
 315 muon (anti-)neutrino beam is generated via $\pi^+ \rightarrow \mu^+ + \nu_\mu$ or $\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$.
 316 The decay of muons and kaons results in an irreducible intrinsic electron neutrino
 317 background. In T2K, this background contamination is $O(< 1\%)$ [39]. There is
 318 also an approximately $\sim 5\%$ “wrong-sign” neutrino background of $\bar{\nu}_\mu$ generated
 319 via the same decays. As the beam is generated by proton interactions (rather
 320 than anti-proton interactions), the wrong-sign component in the antineutrino
 321 beam is larger when operating in neutrino mode.

322 Tuning the proton energy in the beam and using beam focusing techniques
 323 allows the neutrino energy to be set to a value that maximises the disappear-
 324 ance oscillation probability in the L/E term in Equation 2.10. This means that
 325 accelerator experiments are typically more sensitive to the mixing parameters as
 326 compared to a natural neutrino source. However, the disadvantage compared
 327 to atmospheric neutrino experiments is the cost of building a facility to provide

328 high-energy neutrinos, with a high flux, which is required for longer baselines.
 329 Consequently, there is typically less sensitivity to matter effects and the ordering
 330 of the neutrino mass eigenstates.

331 A neutrino experiment measures

$$R(\vec{x}) = \Phi(E_\nu) \times \sigma(E_\nu) \times \epsilon(\vec{x}) \times P(\nu_\alpha \rightarrow \nu_\beta), \quad (2.14)$$

332 where $R(\vec{x})$ is the event rate of neutrinos at position \vec{x} , $\Phi(E_\nu)$ is the flux of
 333 neutrinos with energy E_ν , $\sigma(E_\nu)$ is the cross-section of the neutrino interaction and
 334 $\epsilon(\vec{x})$ is the efficiency and resolution of the detector. In order to leverage the most
 335 out of an accelerator neutrino experiment, the flux and cross-section systematics
 336 need to be constrained. This is typically done via the use of a “near detector”,
 337 situated at a baseline of $O(1)$ km. This detector observes the unoscillated neutrino
 338 flux and constrains the parameters used within the flux and cross-section model.

339 The first accelerator experiments to precisely measure oscillation parameters
 340 were MINOS [40] and K2K [41]. These experiments confirmed the ν_μ disappear-
 341 ance seen in atmospheric neutrino experiments by finding consistent parameter
 342 values for $\sin^2(\theta_{23})$ and Δm_{32}^2 . The current generation of accelerator neutrino
 343 experiments, T2K and NO ν A extended this field by observing $\bar{\nu}_\mu \rightarrow \bar{\nu}_e$ and lead
 344 the sensitivity to atmospheric mixing parameters as seen in Figure 2.6 [42]. The
 345 two experiments differ in their peak neutrino energy, baseline, and detection tech-
 346 nique. The NO ν A experiment is situated at a baseline of 810km from the NuMI
 347 beamline which delivers 2GeV neutrinos. The T2K neutrino beam is peaked
 348 around 0.6GeV and propagates 295km [43]. Additionally, the NO ν A experiment
 349 uses functionally identical detectors (near and far) whereas T2K uses a plastic
 350 scintillator technique at the near detector and a water Cherenkov far detector.
 351 The future generation experiments DUNE [44] and Hyper-Kamiokande [45]
 352 will succeed these experiments as the high-precision era of neutrino oscillation
 353 parameter measurements develops.

354 Several anomalous results have been observed in the LSND [9] and Mini-
 355 BooNE [10] detectors which were designed with purposefully short baselines.

356 Parts of the neutrino community attributed these results to oscillations induced
357 by a fourth “sterile” neutrino [46] but several searches in other experiments,
358 MicroBooNE [47] and KARMEN [48], found no hints of additional neutrino
359 species. The solution to the anomalous results is still being determined.

360 2.3.3 Atmospheric Neutrinos

361 The interactions of primary cosmic ray protons in the Earth’s upper atmosphere
362 generate showers of energetic hadrons. These are mostly pions and kaons that
363 decay to produce a natural source of neutrinos spanning energies of MeV to
364 TeV [49]. The main decay is via,

$$\begin{aligned} \pi^\pm &\rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu), \\ \mu^\pm &\rightarrow e^\pm + (\nu_e, \bar{\nu}_e) + (\nu_\mu, \bar{\nu}_\mu), \end{aligned} \tag{2.15}$$

365 such that for a single pion decay, three neutrinos can be produced. The
366 atmospheric neutrino flux energy spectra as predicted by the Bartol [50], Honda
367 [51–53], and FLUKA [54] models are illustrated in Figure 2.3. The flux distribution
368 peaks at an energy of $O(10)$ GeV. The uncertainties associated with these models
369 are dominated by the hadronic production of kaon and pions as well as the
370 primary cosmic flux.

371 Unlike long-baseline experiments which have a fixed baseline, the distance
372 atmospheric neutrinos propagate is dependent upon the zenith angle at which
373 they interact. This is illustrated in Figure 2.4. Neutrinos that are generated
374 directly above the detector ($\cos(\theta) = 1.0$) have a baseline equivalent to the
375 height of the atmosphere, whereas neutrinos that interact directly below the
376 detector ($\cos(\theta) = -1.0$) have to travel a length equal to the diameter of the Earth.
377 This means atmospheric neutrinos have a baseline that varies from $O(20)$ km to
378 $O(6 \times 10^3)$ km. Any neutrino generated at or below the horizon will be subject
379 to MSW matter resonance as they propagate through the Earth.

380 Figure 2.5 highlights the neutrino flux as a function of the zenith angle for
381 different slices of neutrino energy. For medium to high-energy neutrinos (and to



Figure 2.3: Left panel: The atmospheric neutrino flux for different neutrino flavours as a function of neutrino energy as predicted by the 2007 Honda model (“This work”) [51], the 2004 Honda model (“HKKM04”)[52], the Bartol model [50] and the FLUKA model [54]. Right panel: The ratio of the muon to electron neutrino flux as predicted by all the quoted models. Both figures taken from [51].



Figure 2.4: A diagram illustrating the definition of zenith angle as used in the Super Kamiokande experiment [55].

382 a lesser degree for low-energy neutrinos), the flux is approximately symmetric
 383 around $\cos(\theta) = 0$. To the accuracy of this approximation, the systematic
 384 uncertainties associated with atmospheric flux for comparing upward-going
 385 and down-going neutrino cancels. This allows the down-going events, which are

- 386 mostly insensitive to oscillation probabilities, to act as an unoscillated prediction
 387 (similar to a near detector in an accelerator neutrino experiment).



Figure 2.5: Prediction of ν_e , $\bar{\nu}_e$, ν_μ , $\bar{\nu}_\mu$ fluxes as a function of zenith angle as calculated by the HKKM model [53]. The left, middle and right panels represent three values of neutrino energy, 0.32GeV, 1.0GeV and 3.2GeV respectively. Predictions for other models including Bartol [50], Honda [51] and FLUKA [54] are given in [55].

388 Precursory hints of atmospheric neutrinos were observed in the mid-1960s
 389 searching for $\nu_\mu + X \rightarrow X^* + \mu^\pm$ [56]. This was succeeded by the IMB-3 [57]
 390 and Kamiokande [58] experiments which measured the double ratio of muon
 391 to electron neutrinos in data to Monte Carlo, $R(\nu_\mu/\nu_e) = (\mu/e)_{Data}/(\mu/e)_{MC}$.
 392 Both experiments were found to have a consistent deficit of muon neutrinos,
 393 with $R(\nu_\mu/\nu_e) = 0.67 \pm 0.17$ and $R(\nu_\mu/\nu_e) = 0.658 \pm 0.016 \pm 0.035$, respectively.
 394 Super-Kamiokande (SK) [55] extended this analysis by fitting oscillation pa-
 395 rameters in $P(\nu_\mu \rightarrow \nu_\tau)$ which found best fit parameters $\sin^2(2\theta) > 0.92$ and
 396 $1.5 \times 10^{-3} < \Delta m^2 < 3.4 \times 10^{-3} eV^2$.

397 Since then, atmospheric neutrino experiments have been making precision
 398 measurements of the $\sin^2(\theta_{23})$ and Δm^2_{32} oscillation parameters. Atmospheric
 399 neutrino oscillation is dominated by $P(\nu_\mu \rightarrow \nu_\tau)$, where SK observed a 4.6σ
 400 discovery of ν_τ appearance [59]. Figure 2.6 illustrates the current estimates on
 401 the atmospheric mixing parameters, from a wide range of atmospheric and
 402 accelerator neutrino observatories.



Figure 2.6: Constraints on the atmospheric oscillation parameters, $\sin^2(\theta_{23})$ and Δm_{32}^2 , from atmospheric and long-baseline experiments: SK [60], T2K [61], NOvA [62], IceCube [63] and MINOS [64]. Figure taken from [65].

2.3.4 Reactor Neutrinos

As illustrated in the first discovery of neutrinos (section 2.1), nuclear reactors are a very useful artificial source of electron antineutrinos. For reactors that use low-enriched uranium ^{235}U as fuel, the antineutrino flux is dominated by the β -decay fission of ^{235}U , ^{238}U , ^{239}Pu and ^{241}Pu [66] as illustrated in Figure 2.7.

Due to their low energy, reactor electron antineutrinos predominantly interact via the inverse β -decay (IBD) interaction. The typical signature contains two signals delayed by $O(200)\mu\text{s}$; firstly the prompt photons from positron annihilation, and secondly the photon emitted ($E_{tot}^\gamma = 2.2\text{MeV}$) from de-excitation after neutron capture on hydrogen. Searching for both signals improves the detector's ability to distinguish between background and signal events [68].

There are many short baseline experiments ($L \sim O(1)\text{km}$) that have measured the $\sin^2(\theta_{13})$ and Δm_{32}^2 oscillation parameters. Daya Bay [69], RENO [70] and Double Chooz [71] have all provided precise measurements, with the first discovery of a non-zero θ_{13} made by Daya Bay and RENO (and complemented by T2K [71]). The constraints on $\sin^2(\theta_{13})$ by the reactor experiments lead the field. They



Figure 2.7: Reactor electron antineutrino fluxes for ^{235}U (Black), ^{238}U (Green), ^{239}Pu (Purple), and ^{241}Pu (Orange) isotopes. The inverse β -decay cross-section (Blue) and corresponding measurable neutrino spectrum (Red) are also given. Top panel: Schematic of Inverse β -decay interaction including the eventual capture of the emitted neutron. This capture emits a γ -ray which provides a second signal of the event. Taken from [67].

are often used as external inputs to accelerator neutrino experiments to improve their sensitivity to δ_{CP} and mass hierarchy determination. JUNO-TAO [72], a small collaboration within the larger JUNO experiment, is a next-generation reactor experiment that aims to precisely measure the isotopic antineutrino yields from the different fission chains.

Kamland [73] is the only experiment to have observed reactor neutrinos using a long baseline (flux weighted averaged baseline of $L \sim 180\text{km}$) which allows it to have sensitivity to Δm_{21}^2 . Combined with the SK solar neutrino experiment, the combined analysis puts the most stringent constraint on Δm_{21}^2 [74].

2.4 Summary Of Oscillation Parameter Measurements

Since the first evidence of neutrino oscillations, numerous measurements of the mixing parameters have been made. Many experiments use neutrinos as a tool for the discovery of new physics (diffuse supernova background, neutrinoless double beta decay and others) so the PMNS parameters are summarised in the Particle Data Group (PDG) review tables. The analysis presented in this thesis focuses on the 2020 T2K oscillation analysis presented in [75] which the 2020 PDG constraints [76] were used. These constraints are outlined in Table 2.1.

Parameter	2020 Constraint
$\sin^2(\theta_{12})$	0.307 ± 0.013
Δm_{21}^2	$(7.53 \pm 0.18) \times 10^{-5} \text{ eV}^2$
$\sin^2(\theta_{13})$	$(2.18 \pm 0.07) \times 10^{-2}$
$\sin^2(\theta_{23})$ (I.H.)	0.547 ± 0.021
$\sin^2(\theta_{23})$ (N.H.)	0.545 ± 0.021
Δm_{32}^2 (I.H.)	$(-2.546^{+0.034}_{-0.040}) \times 10^{-3} \text{ eV}^2$
Δm_{32}^2 (N.H.)	$(2.453 \pm 0.034) \times 10^{-3} \text{ eV}^2$

Table 2.1: The 2020 Particle Data Group constraints of the oscillation parameters taken from [76]. The value of Δm_{32}^2 is given for both normal hierarchy (N.H.) and inverted hierarchy (I.H.) and $\sin^2(\theta_{23})$ is broken down by whether its value is below (Q1) or above (Q2) 0.5.

The $\sin^2(\theta_{13})$ measurement stems from the electron antineutrino disappearance, $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$, and is taken as the average best-fit from the combination of Daya Bay, Reno and Double Chooz. It is often used as a prior uncertainty within other neutrino oscillation experiments, typically termed the reactor constraint. The $\sin^2(\theta_{12})$ parameter is predominantly measured through electron neutrino disappearance, $P(\nu_e \rightarrow \nu_{\mu,\tau})$, in solar neutrino experiments. The long-baseline reactor neutrino experiment Kamland also has a sensitivity to this parameter and is used in a joint fit to solar data from SNO and SK, using the reactor constraint. Measurements of $\sin^2(\theta_{23})$ are made by long-baseline and atmospheric neutrino experiments. The PDG value is a joint fit of T2K, NO ν A, MINOS and IceCube DeepCore experiments. The latest T2K-only measurement, provided at Neutrino2020 and is the basis of this thesis, is given as $\sin^2(\theta_{23}) = 0.546^{+0.024}_{-0.046}$ [75].

The PDG constraint on Δm_{21}^2 is provided by the KamLAND experiment using solar and geoneutrino data. This measurement utilised a $\sin^2(\theta_{13})$ constraint from accelerator (T2K, MINOS) and reactor neutrino (Daya Bay, RENO, Double Chooz) experiments. Accelerator measurements make some of the most stringent constraints on Δm_{32}^2 although atmospheric experiments have more sensitivity to the mass hierarchy determination. The PDG performs a joint fit of accelerator and atmospheric data, in both normal and inverted hierarchies separately. The latest T2K-only result is $\Delta m_{32}^2 = 2.49^{+0.058}_{-0.082} \times 10^{-3} \text{ eV}^2$ favouring normal hierarchy [75]. The value of δ_{CP} is largely undetermined. CP-conserving values of 0 and π were rejected with $\sim 2\sigma$ intervals, as published in Nature, although more recent analyses have reduced the credible intervals to 90%. Since the 2020 PDG publication, there has been a new measurement of $\sin^2(\theta_{13}) = (2.20 \pm 0.07) \times 10^{-2}$ [77], alongside updated Δm_{32}^2 and $\sin^2(\theta_{23})$ measurements.

Throughout this thesis, several sample spectra predictions and contours are presented, which require oscillation parameters to be assumed. Table 2.2 defines two sets of oscillation parameters, with “Asimov A” set being close to the preferred values from a previous T2K-only fit [78] and “Asimov B” being CP-conserving and further from maximal θ_{23} mixing.

Parameter	Asimov A	Asimov B
Δm_{12}^2	$7.53 \times 10^{-5} \text{ eV}^2$	
Δm_{32}^2	$2.509 \times 10^{-3} \text{ eV}^2$	
$\sin^2(\theta_{12})$	0.304	
$\sin^2(\theta_{13})$	0.0219	
$\sin^2(\theta_{23})$	0.528	0.45
δ_{CP}	-1.601	0.0

Table 2.2: Reference values of the neutrino oscillation parameters for two different oscillation parameter sets.

2.5 Overview of Oscillation Effects

The analysis presented within this thesis focuses on the determination of oscillation parameters from atmospheric and beam neutrinos. Whilst subject to the

469 same oscillation formalism, the way in which the two samples have sensitivity
470 to the different oscillation parameters differs significantly.

471 Atmospheric neutrinos have a varying baseline, or “path length” L , such that
472 the distance each neutrino travels before interacting is dependent upon the zenith
473 angle, θ_Z . As primary cosmic rays can interact anywhere between the Earth’s
474 surface and $\sim 50\text{km}$ above that, the height, h , in the atmosphere at which the
475 neutrino was generated also affects the path length,

$$L = \sqrt{(R_E + h)^2 - R_E^2(1 - \cos^2(\theta_Z))} - R_E \cos(\theta_Z). \quad (2.16)$$

476 Where $R_E = 6,371\text{km}$ is the Earth’s radius. This assumes a spherically
477 symmetric Earth model. Therefore, the oscillation probability is dependent upon
478 two parameters, $\cos(\theta_Z)$ and E_ν .

479 The oscillation probability used within this analysis is based on [21]. The
480 neutrino wavefunction in the vacuum Hamiltonian evolves in each layer of
481 constant matter density via

$$i \frac{d\psi_j(t)}{dt} = \frac{m_j^2}{2E_\nu} \psi_j(t) - \sum_k \sqrt{2} G_F N_e U_{ej} U_{ke}^\dagger \psi_k(t), \quad (2.17)$$

482 where m_j^2 is the square of the j^{th} vacuum eigenstate mass, E_ν is the neutrino
483 energy, G_F is Fermi’s constant, N_e is the electron number density and U is the
484 PMNS matrix. The transformation $N_e \rightarrow -N_e$ and $\delta_{CP} \rightarrow -\delta_{CP}$ is applied for
485 antineutrino propagation. Thus, a model of the Earth’s density is required for
486 neutrino propagation. Following the official SK-only methodology [79], this
487 analysis uses the Preliminary Reference Earth Model (PREM) [80] which provides
488 piecewise cubic polynomials as a function of the Earth’s radius. This density
489 profile is illustrated in Figure 2.8. As the propagator requires layers of constant
490 density, the SK methodology approximates the PREM model by using four layers
491 of constant density [79], detailed in Table 2.3.

492 The atmospheric neutrino oscillation probabilities can be presented as two di-
493 mensional “oscillograms” as illustrated in Figure 2.9. The distinct discontinuities,
494 as a function of $\cos(\theta_Z)$, are due to the discontinuous density in the PREM model.

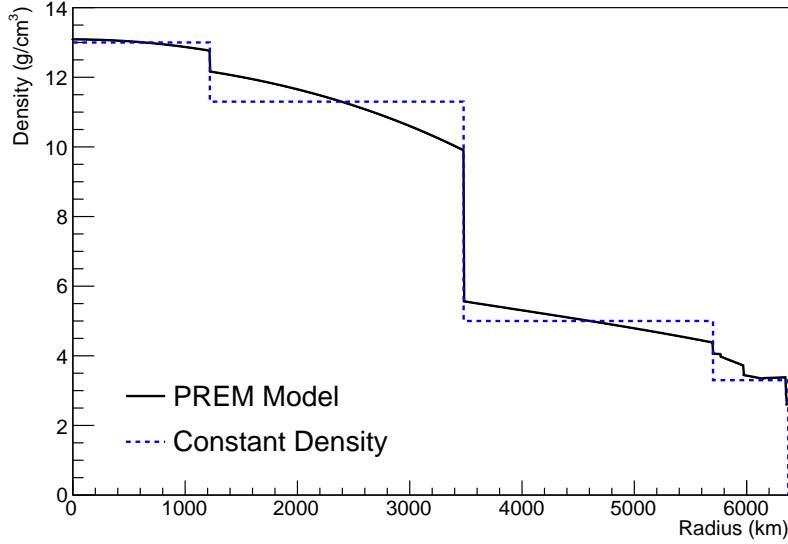


Figure 2.8: The density of the Earth given as a function of the radius, as given by the PREM model (Black), and the constant density four-layer approximation (Blue), as used in the official SK-only analysis.

Layer	Outer Radius [km]	Density [g/cm ³]	Chemical composition (Z/A)
Inner Core	1220	13	0.468 ± 0.029
Outer Core	3480	11.3	0.468 ± 0.029
Lower Mantle	5701	5.0	0.496
Transition Zone	6371	3.3	0.496

Table 2.3: Description of the four layers of the Earth invoked within the constant density approximation of the PREM model [80].

Atmospheric neutrinos have sensitivity to δ_{CP} through the overall event rate. Figure 2.10 illustrates the difference in oscillation probability between CP-conserving ($\delta_{CP} = 0.$) and a CP-violating ($\delta_{CP} = -1.601$) value taken from Asimov A oscillation parameter set (Table 2.2). The result is a complicated oscillation pattern in the appearance probability for sub-GeV upgoing neutrinos. The detector does not have sufficient resolution to resolve these individual patterns so the sensitivity to δ_{CP} for atmospheric neutrinos comes via the overall normalisation of these events.

The presence of matter means that the effect δ_{CP} has on the oscillation probability is not equal between neutrinos and antineutrinos. Furthermore, the interaction cross-section for neutrinos is larger than for antineutrinos so the two

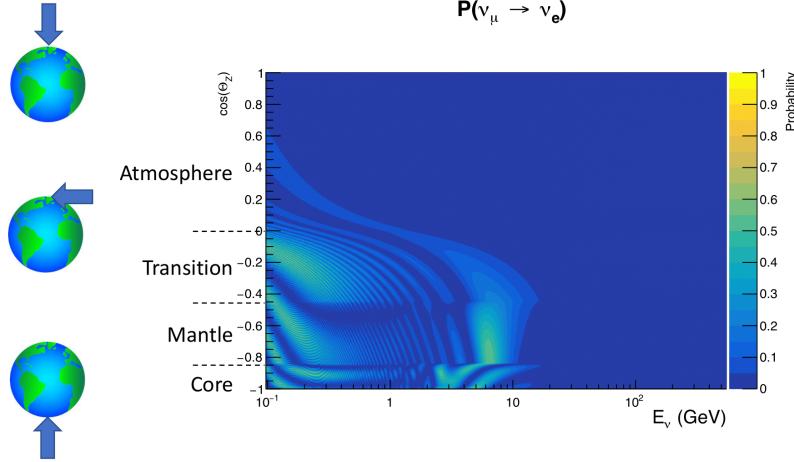


Figure 2.9: An “oscillogram” that depicts the $P(\nu_\mu \rightarrow \nu_e)$ oscillation probability as a function of neutrino energy and cosine of the zenith angle. The zenith angle is defined such that $\cos(\theta_Z) = 1.0$ represents neutrinos that travel from directly above the detector. The four-layer constant density PREM model approximation is used and Asimov A oscillation parameters are assumed (Table 2.2).

506 effects have to be disentangled. These effects are further convoluted by detector
 507 efficiencies as SK cannot distinguish neutrinos and antineutrinos well. All of
 508 these effects lead to a difference in the number of neutrinos detected compared
 509 to antineutrinos. This changes how the δ_{CP} normalisation term is observed,
 510 resulting in a very complex sensitivity to δ_{CP} .

511 The vacuum and matter oscillation probabilities for $P(\nu_e \rightarrow \nu_e)$ and $P(\bar{\nu}_e \rightarrow$
 512 $\bar{\nu}_e)$ are presented in Figure 2.11, where the PREM model has been assumed. The
 513 oscillation probability for both neutrinos and antineutrinos is affected in the
 514 presence of matter. However, the resonance effects around $O(5)\text{GeV}$ only occur
 515 for neutrinos in the normal mass hierarchy and antineutrinos in the inverse mass
 516 hierarchy. The exact position and amplitude of the resonance depend on $\sin^2(\theta_{23})$,
 517 further increasing the atmospheric neutrinos’ sensitivity to the parameter.

518 As the T2K beam flux is centered at the first oscillation maximum ($E_\nu =$
 519 0.6GeV) [43], the sensitivity to δ_{CP} is predominantly observed as a change in the
 520 event-rate of e-like samples in $\nu/\bar{\nu}$ modes. Figure 2.12 illustrates the $P(\nu_\mu \rightarrow \nu_e)$
 521 oscillation probability for a range of δ_{CP} values. A circular modulation of the

$$\mathbf{P}(\nu_\mu \rightarrow \nu_e; \delta_{CP} = -1.601) - \mathbf{P}(\nu_\mu \rightarrow \nu_e; \delta_{CP} = 0.)$$

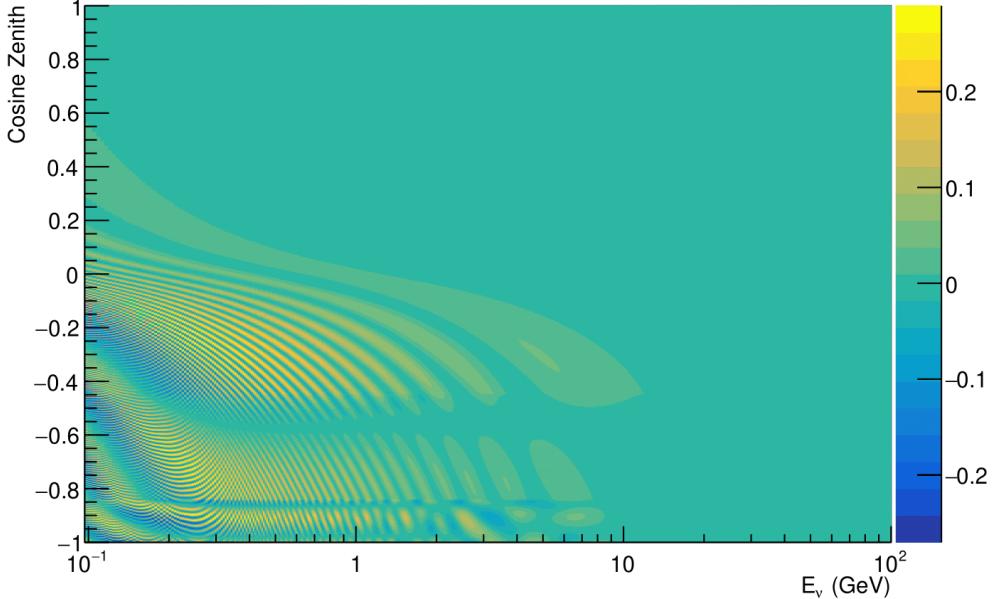


Figure 2.10: The effect of δ_{CP} for atmospheric neutrinos given in terms of the neutrino energy and zenith angle. This oscillogram compares the $P(\nu_\mu \rightarrow \nu_e)$ oscillation probability for a CP conserving ($\delta_{CP} = 0.0$) and a CP violating ($\delta_{CP} = -1.601$) value taken from the Asimov A parameter set. The other oscillation parameters assume the Asimov A oscillation parameter set given in Table 2.2.

522 first oscillation peak (in both magnitude and position) is observed when varying
 523 throughout the allowable values of δ_{CP} . The CP-conserving values of $\delta_{CP} = 0, \pi$
 524 have a lower(higher) oscillation maximum than the CP-violating values of $\delta_{CP} =$
 525 $-\pi/2$ ($\delta_{CP} = \pi/2$). A sub-dominant shift in the energy of the oscillation peak is
 526 also present, which aids in separating the two CP-conserving values of δ_{CP} .

527 T2K's sensitivity to $\sin^2(\theta_{23})$ and Δm_{32}^2 is observed as a shape-based variation
 528 of the muon-like samples, as illustrated in Figure 2.12. The value of Δm_{32}^2 laterally
 529 shifts the position of the oscillation dip (around $E_\nu \sim 0.6\text{GeV}$) in the $P(\nu_\mu \rightarrow \nu_\mu)$
 530 oscillation probability. A variation of $\sin^2(\theta_{23})$ is predominantly observed as
 531 a vertical shift of the oscillation dip with second-order horizontal shifts being
 532 due to matter effects. The beam neutrinos have limited sensitivity to matter
 533 effects due to the relatively shorter baseline as well as the Earth's mantle being
 534 a relatively low-density material (as compared to the Earth's core). For some



Figure 2.11: An illustration of the matter-induced effects on the oscillation probability, given as a function of neutrino energy and zenith angle. The top row of panels gives the $P(\nu_e \rightarrow \nu_e)$ oscillation probability and the bottom row illustrates the $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$ oscillation probability. The left column highlights the oscillation probability in a vacuum, whereas the middle and right column represents the oscillation probabilities when the four-layer fixed density PREM model is assumed. All oscillation probabilities assume the “Asimov A” set given in Table 2.2, but importantly, the right column sets an inverted mass hierarchy. The “matter resonance” effects at $E_\nu \sim 5\text{GeV}$ can be seen in the $P(\nu_e \rightarrow \nu_e)$ for normal mass hierarchy and $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$ for inverted hierarchy.

535 values of δ_{CP} , the degeneracy in the number of e-like events allows the mass
 536 hierarchy to be broken. This leads to a δ_{CP} -dependent mass hierarchy sensitivity
 537 which can be seen in Figure 2.13.

538 Whilst all oscillation channels should be included for completeness, the
 539 computational resources required to run a fit are limited and any reasonable
 540 approximations which reduce the number of oscillation probability calculations
 541 that need to be made should be applied. The $\nu_e \rightarrow \nu_{e,\mu,\tau}$ (and antineutrino

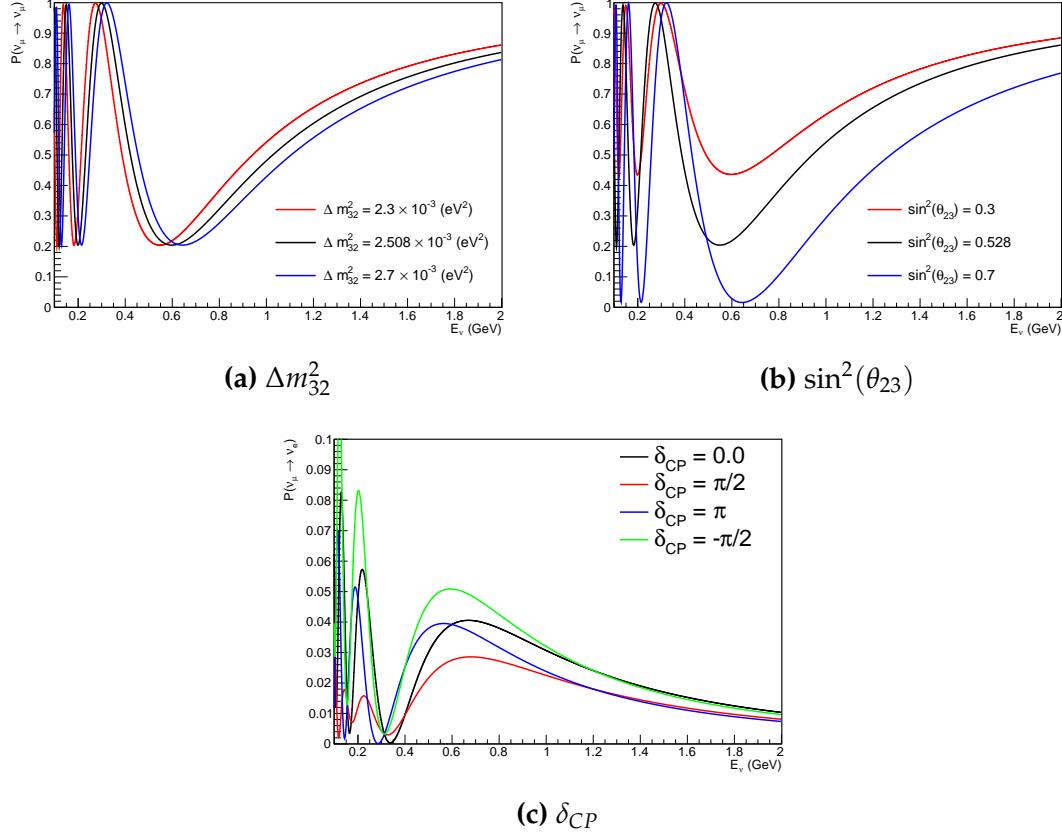


Figure 2.12: The oscillation probability for beam neutrino events given as a function of neutrino energy. All oscillation parameters assume the “Asimov A” set given in Table 2.2 unless otherwise stated. Each panel represents a change in one of the oscillation parameters whilst keeping the remaining parameters fixed.

equivalent) oscillations can be ignored for beam neutrinos as the $\nu_e/\bar{\nu}_e$ fluxes are approximately two orders of magnitude smaller than the corresponding $\nu_\mu/\bar{\nu}_\mu$ flux. Furthermore, as the peak neutrino energy of the beam is well below the threshold for charged current tau production ($E_\nu = 3.5\text{GeV}$ [59]), only a small proportion of the neutrinos produced in the beam have the required energy. For the few neutrinos that have sufficient energy, the oscillation probability is very small due to their energy being well above the oscillation maximum (small value of L/E). Whilst these approximations have been made for the beam neutrinos, the atmospheric flux of ν_e is of the same order of magnitude as the ν_μ flux and the energy distribution of atmospheric neutrinos extends well above the tau production threshold. These events can have non-negligible oscillation



Figure 2.13: The number of electron-like events in the FHC and RHC operating mode of the beam, as a function of the oscillation probabilities. Both normal hierarchy (Solid) and inverse hierarchy (Dashed) values of Δm_{32}^2 are given.

⁵⁵³ probabilities due to the further distance they travel.

3

554

555

T2K and SK Experiment Overview

556 As the successor of the Kamiokande experiment, the Super-Kamiokande (SK)
557 collaboration has been leading atmospheric neutrino oscillation analyses for
558 over two decades. The detector has provided some of the strongest constraints
559 on proton decay and the first precise measurements of the Δm_{32}^2 and $\sin^2(\theta_{23})$
560 neutrino oscillation parameters. The history, detection technique, and operation
561 of the SK detector is described in section 3.1.

562 The Tokai-to-Kamioka (T2K) experiment was one of the first long-baseline
563 experiments to use both neutrino and antineutrino beams to precisely measure
564 charge parity violation within the neutrino sector. The T2K experiment observed
565 the first hints of a non-zero $\sin^2(\theta_{13})$ measurement and continues to lead the
566 field with the constraints it provides on $\sin^2(\theta_{13})$, $\sin^2(\theta_{23})$, Δm_{32}^2 and δ_{CP} . In
567 section 3.2, the techniques that T2K use to generate the neutrino beam and
568 constrain systematic parameter through near detector constraints are described.

569 3.1 The Super-Kamiokande Experiment

570 The SK experiment began taking data in 1996 [81] and has had many modifi-
571 cations throughout its operation. There have been seven defined periods of
572 data taking as noted in Table 3.1. Data taking began in SK-I which ran for five

years. Between the SK-I and SK-II periods, approximately 55% of the PMTs were damaged during maintenance [82]. Those that survived were equally distributed throughout the detector in the SK-II era, which resulted in a reduced 19% photo-coverage. From SK-III onwards, repairs to the detector meant the full suite of PMTs was operational recovering the 40% photo-coverage. Before the start of SK-IV, the data acquisition and electronic systems were upgraded. Between SK-IV and SK-V, a significant effort was placed into tank open maintenance and repair/replacement of defective PMTs in preparation for the Gadolinium upgrade; a task for which the author of this thesis was required. Consequently, the detector conditions were significantly changed from this point. SK-VI marked the start of the SK-Gd era, with the detector being doped with gadolinium at a concentration of 0.01% by concentration. SK-VII, which started during the writing of this thesis, has increased the gadolinium concentration to 0.03% for continued operation [83].

The oscillation analysis presented within this thesis focuses on the SK-IV period of running and the data taken within it. This follows from the recent SK analysis presented in [84]. Therefore, the information presented within this section focuses on that period.

Period	Start Date	End Date	Live-time (days)
I	April 1996	July 2001	1489.19
II	October 2002	October 2005	798.59
III	July 2006	September 2008	518.08
IV	September 2008	May 2018	3244.4
V	January 2019	July 2020	461.02
VI	July 2020	May 2022	583.3
VII	May 2022	Ongoing	N/A

Table 3.1: The various SK periods and their respective live-time. The SK-VI live-time is calculated until 1st April 2022. SK-VII started during the writing of this thesis.

3.1.1 The SK Detector

The basic structure of the Super-Kamiokande (SK) detector is a cylindrical tank with a diameter 39.3m and height 41.1m filled with ultrapure water [82]. A diagram of the significant components of the SK detector is given in Figure 3.1.

594 The SK detector is situated in the Kamioka mine in Gifu, Japan. The mine
 595 is underground with roughly 1km rock overburden (2.7km water equivalent
 596 overburden) [85]. At this depth, the rate of cosmic ray muons is significantly
 597 decreased to a value of $\sim 2\text{Hz}$ (net rate) [86]. The top of the tank is covered
 598 with stainless steel which is designed as a working platform for maintenance,
 599 calibration, and location for high voltage and data acquisition electronics.



Figure 3.1: A schematic diagram of the Super-Kamiokande Detector. Taken from [87].

600 A smaller cylindrical structure (36.2m diameter, 33.8m height) is situated
 601 inside the tank, with an approximate 2m gap between this structure and the outer
 602 tank wall. The purpose of this structure is to support the photomultiplier tubes
 603 (PMTs). The volume inside and outside the support structure is referred to as the
 604 inner detector (ID) and outer detector (OD), respectively. In the SK-IV era, the
 605 ID and OD are instrumented by 11,129 50cm and 1,885 20cm PMTs respectively
 606 [82]. The ID contains a 32kton mass of water. Many analyses performed at SK
 607 use a “fiducial volume” defined by the volume of water inside the ID excluding
 608 some distance to the ID wall. This reduces the volume of the detector which is
 609 sensitive to neutrino events but reduces radioactive backgrounds and allows for

610 better reconstruction performance. The nominal fiducial volume is defined as the
611 area contained inside 2m from the ID wall for a total of 22.5kton water [88].

612 The two regions of the detector (ID and OD) are optically separated with
613 opaque black plastic hung from the support structure. The purpose of this is
614 to determine whether an event entered or exited the ID. This allows cosmic ray
615 muons and partially contained events to be tagged and separated from neutrino
616 events entirely contained within the ID. This black plastic is also used to cover
617 the area between the ID PMTs to reduce photon reflection from the ID walls.
618 Opposite to this, the OD is lined with a reflective material to allow photons to
619 reflect around inside the OD until collected by one of the PMTs. Furthermore,
620 each OD PMT is optically coupled with $50 \times 50\text{cm}$ plates of wavelength shifting
621 acrylic which increases the efficiency of light collection [85].

622 In the SK-IV data-taking period, the photocathode coverage of the detector, or
623 the fraction of the ID wall instrumented with PMTs, is $\sim 40\%$ [85]. The PMTs have
624 a quantum efficiency (the ratio of detected electrons to incident photons) of $\sim 21\%$
625 for photons with wavelengths of $360\text{nm} < \lambda < 390\text{nm}$ [89, 90]. The proportion
626 of photoelectrons that produce a signal in the dynode of a PMT, termed the
627 collection efficiency, is $> 70\%$ [85]. The PMTs used within SK are most sensitive
628 to photons with wavelength $300\text{nm} \leq \lambda \leq 600\text{nm}$ [85]. One disadvantage of
629 using PMTs as the detection media is that the Earth's geomagnetic field can
630 modify its response. Therefore, a set of compensation coils is built around the
631 inner surface of the detector to mitigate this effect [86].

632 The SK detector is filled with ultrapure water, which in a perfect world, con-
633 tains no impurities. However, bacteria and organic compounds can significantly
634 degrade the water quality. This decreases the attenuation length, which reduces
635 the total number of photons that hit a PMT. To combat this, a sophisticated water
636 treatment system has been developed [85, 91]. UV lights, mechanical filters, and
637 membrane degasifiers are used to reduce the bacteria, suspended particulates,
638 and radioactive materials from the water. The flow of water within the tank
639 is also critical as it can remove stagnant bacterial growth or build-up of dust

640 on the surfaces within the tank. Gravity drifts impurities in the water towards
641 the bottom of the tank which, if left uncontrolled, can create asymmetric water
642 conditions between the top and bottom of the tank. Typically, the water entering
643 the tank is cooled below the ambient temperature of the tank to control convection
644 and inhibit bacteria growth. Furthermore, the rate of dark noise hits within PMTs
645 is sensitive to the PMT temperature [92]. Therefore controlling the temperature
646 gradients within the tank is beneficial for stable measurements.

647 SK-VI is the first phase of the SK experiment to use gadolinium dopants
648 within the ultrapure water [83]. As such, the SK water system had to be replaced
649 to avoid removing the gadolinium concentrate from the ultrapure water [93]. For
650 an inverse β -decay (IBD) interaction on a water target, the emitted neutron is
651 thermally captured on hydrogen. This process releases a 2.2MeV γ ray which is
652 difficult to detect as the resulting Compton scattered electrons are very close to the
653 Cherenkov threshold, limiting detection capability. Thermal capture of neutrons
654 on gadolinium generates γ rays with higher energy (8MeV [68]) meaning they
655 are more easily detected and reconstructed. SK-VI has 0.01% Gd loading (0.02%
656 gadolinium sulphate by mass) which causes \approx 50% of neutrons emitted by IBD to
657 be captured on gadolinium[94, 95] . Whilst predominantly useful for low energy
658 analyses, Gd loading allows better $\nu/\bar{\nu}$ separation for atmospheric neutrino
659 event selections [96]. Efforts are currently in place to increase the gadolinium
660 concentrate to 0.03% for \approx 75% neutron capture efficiency on gadolinium [97].
661 The final stage of loading targets 0.1% concentrate for \approx 90% neutron capture
662 efficiency on gadolinium.

663 3.1.2 Calibration

664 The calibration of the SK detector is documented in [82] and summarised below.
665 The analysis presented within this thesis is dependent upon ‘high energy events’
666 (Charged particles with $O(> 100)$ MeV momenta). These are events that are
667 expected to generate a larger number of photons such that each PMT will
668 be hit with multiple photons. The reconstruction of these events depends

upon the charge deposited within each PMT and the timing response of each individual PMT. Therefore, the most relevant calibration techniques to this thesis are outlined.

Before installation, 420 PMTs were calibrated to have identical charge responses and then distributed throughout the tank in a cross-shape pattern (As illustrated by Figure 3.2). These are used as a standardised measure for the rest of the PMTs installed at similar geometric positions within SK to be calibrated against. To perform this calibration, a xenon lamp is located at the center of the SK tank which flashes uniform light at 1Hz. This allows for geometrical effects, water quality variation, and timing effects to be measured in situ throughout normal data-taking periods.

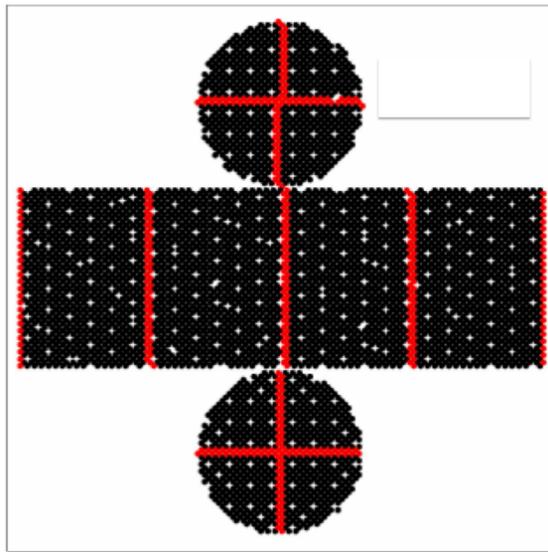


Figure 3.2: The location of “standard PMTs” (red) inside the SK detector. Taken from [82].

When specifically performing calibration of the detector (in out-of-data taking mode), the water in the tank was circulated to avoid top/bottom asymmetric water quality. Any non-uniformity within the tank significantly affects the PMT hit probability through scattering or absorption. This becomes a dominant effect for very low-intensity light sources that are designed such that only one photon is incident upon a given PMT.

686 The gain of a PMT is defined as the ratio of the total charge of the signal
 687 produced compared to the charge of photoelectrons emitted by the photocathodes
 688 within the PMT. To calibrate the signal of each PMT, the “relative” and “absolute”
 689 gain values are measured. The relative gain is the variation of gain among each
 690 of the PMTs whereas the absolute gain is the average gain of all PMTs.

691 The relative gain is calibrated as follows. A laser is used to generate two
 692 measurements: a high-intensity flash that illuminates every PMT with a sufficient
 693 number of photons, and a low-intensity flash in which only a small number
 694 of PMTs collect light. The first measurement creates an average charge, $Q_{obs}(i)$
 695 on PMT i , whereas the second measurement ensures that each hit PMT only
 696 generates a single photoelectron. For the low-intensity measurement, the number
 697 of times each PMT records a charge larger than 1/4 photoelectrons, $N_{obs}(i)$, is
 698 counted. The values measured can be expressed as

$$\begin{aligned} Q_{obs}(i) &\propto I_H \times f(i) \times \epsilon(i) \times G(i), \\ N_{obs}(i) &\propto I_L \times f(i) \times \epsilon(i). \end{aligned} \tag{3.1}$$

699 Where I_H and I_L is the intensity of the high and low flashes, $f(i)$ is the
 700 acceptance efficiency of the i^{th} PMT, $\epsilon(i)$ is the product of the quantum and
 701 collection efficiency of the i^{th} PMT and $G(i)$ is the gain of the i^{th} PMT. The relative
 702 gain for each PMT can be determined by taking the ratio of these quantities.

703 The absolute gain calibration is performed by observing fixed energy γ -rays
 704 of $E_\gamma \sim 9\text{MeV}$ emitted isotropically from neutron capture on a NiCf source
 705 situated at the center of the detector. This generates a photon yield of about 0.004
 706 photoelectrons/PMT/event, meaning that $> 99\%$ of PMT signals are generated
 707 from single photoelectrons. A charge distribution is generated by performing
 708 this calibration over all PMTs, and the average value of this distribution is taken
 709 to be the absolute gain value.

710 As mentioned in subsection 3.1.1, the average quantum and collection effi-
 711 ciency for the SK detector PMTs is $\sim 21\%$ and $> 70\%$ respectively. However,
 712 these values do differ between each PMT and need to be calibrated accordingly.

713 Consequently, the NiCf source is also used to calibrate the “quantum \times collection”
714 efficiency (denoted “QE”) value of each PMT. The NiCf low-intensity source is
715 used as the PMT hit probability is proportional to the QE ($N_{obs}(i) \propto \epsilon(i)$ in
716 Equation 3.1). A Monte Carlo prediction which includes photon absorption,
717 scattering, and reflection is made to estimate the number of photons incident on
718 each PMT and the ratio of the number of predicted to observed hits is calculated.
719 The difference is attributed to the QE efficiency of that PMT. This technique is
720 extended to calculate the relative QE efficiency by normalizing the average of
721 all PMTs which removes the dependence on the light intensity.

722 Due to differing cable lengths and readout electronics, the timing response
723 between a photon hitting the PMT and the signal being captured by the data
724 acquisition can be different between each PMT. Due to threshold triggers (De-
725 scribed in subsection 3.1.3), the time at which a pulse reaches a threshold is
726 dependent upon the size of the pulse. This is known as the ‘time-walk’ effect
727 and also needs to be accounted for in each PMT. To calibrate the timing response,
728 a pulse of light with width 0.2ns is emitted into the detector through a diffuser.
729 Two-dimensional distributions of time and pulse height (or charge) are made
730 for each PMT and are used to calibrate the timing response. This is performed
731 in-situ during data taking with the light source pulsing at 0.03Hz.

732 The top/bottom water quality asymmetry is measured using the NiCf calibra-
733 tion data and cross-referencing these results to the “standard PMTs”. The water
734 attenuation length is continuously measured by the rate of vertically-downgoing
735 cosmic-ray muons which enter via the top of the tank.

736 Dark noise is where a PMT registers a pulse that is consistent with a single
737 photoelectron emitted from photon detection despite the PMT being in complete
738 darkness. This is predominately caused by two processes. Firstly there is
739 intrinsic dark noise which is where photoelectrons gain enough thermal energy
740 to be emitted from the photocathode, and secondly, the radioactive decay of
741 contaminants inside the structure of the PMT. Typical dark noise rate for PMTs
742 used within SK are $O(3)\text{kHz}$ [85]. This is lower than the expected number of

743 photons generated for a ‘high energy event’ (As described in subsection 3.1.4)
744 but instability in this value can cause biases in reconstruction. Dark noise is
745 related to the gain of a PMT and is calibrated using hits inside a time window
746 recorded before an event trigger [98].

747 3.1.3 Data Acquisition and Triggering

748 As the analysis presented in this thesis will only use the SK-IV period of the
749 SK experiment so this subsection focuses on the relevant points of the data
750 acquisition and triggering systems to that SK period. The earlier data acquisition
751 and triggering systems are documented in [99, 100].

752 Before the SK-IV period started, the existing front-end electronics were re-
753 placed with “QTC-Based Electronics with Ethernet, QBEE” systems [101]. When
754 the QBEE observes a signal above a 1/4 photoelectron threshold, the charge-to-
755 time (QTC) converter generates a rectangular pulse. The start of the rectangular
756 pulse indicates the time at which the analog photoelectron signal was received
757 and the width of the pulse indicates the total charge integrated throughout the
758 signal. This is then digitized by time-to-digital converters and sent to the “front-
759 end” PCs. The digitized signal from every QBEE is then chronologically ordered
760 and sent to the “merger” PCs. It is the merger PCs that apply the software trigger.
761 Any triggered events are passed to the “organizer” PC. This sorts the data stream
762 of multiple merger PCs into chronologically ordered events, which are then saved
763 to disk. The schematic of data flow from PMTs to disk is illustrated in Figure 3.3.

764 The software trigger (described in [103]) operates by determining the number
765 of PMT hits within a 200ns sliding window, N_{200} . This window coincides with
766 the maximum time that a Cherenkov photon would take to traverse the length
767 of the SK tank [100]. For lower energy events that generate fewer photons, this
768 technique is useful for eliminating background processes like dark noise and
769 radioactive decay which would be expected to be separated in time. When the
770 value of N_{200} exceeds some pre-defined threshold, a software trigger is issued.
771 There are several trigger thresholds used within the SK-IV period which are

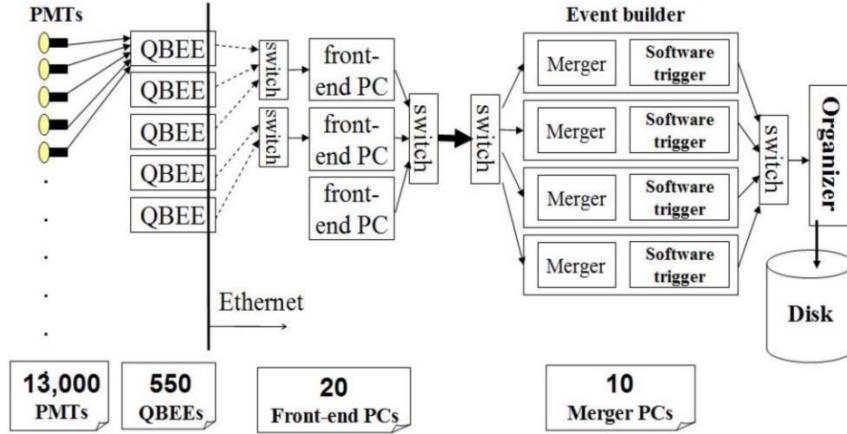


Figure 3.3: Schematic view of the data flow through the data acquisition and online system. Taken from [102].

772 detailed in Table 3.2. If one of these thresholds is met, the PMT hits within an
 773 extended time window are also read out and saved to disk. In the special case
 774 of an event that exceeds the SHE trigger but does not exceed the OD trigger,
 775 the AFT trigger looks for delayed coincidences of 2.2MeV gamma rays emitted
 776 from neutron capture in a $535\mu\text{s}$ window after the SHE trigger. A similar but
 777 more complex “Wideband Intelligent Trigger (WIT)” has been deployed and
 778 is described in [104].

Trigger	Acronym	Condition	Extended time window (μs)
Super Low Energy	SLE	>34/31 hits	1.3
Low Energy	LE	>47 hits	40
High Energy	HE	>50 hits	40
Super High Energy	SHE	>70/58 hits	40
Outer Detector	OD	>22 hits in OD	N/A

Table 3.2: The trigger thresholds and extended time windows saved around an event which were utilised throughout the SK-IV period. The exact thresholds can change and the values listed here represent the thresholds at the start and end of the SK-IV period.

779 3.1.4 Cherenkov Radiation

780 Cherenkov light is emitted from any highly energetic charged particle traveling
 781 with relativistic velocity, β , greater than the local speed of light in a medium [105].

782 Cherenkov light is formed at the surface of a cone with a characteristic pitch angle,

$$\cos(\theta) = \frac{1}{\beta n}. \quad (3.2)$$

783 Where n is the refractive index of the medium. Consequently, the Cherenkov
 784 momentum threshold, P_{thres} , is dependent upon the mass, m , of the charged
 785 particle moving through the medium,

$$P_{thres} = \frac{m}{\sqrt{n^2 - 1}}. \quad (3.3)$$

786 For water, where $n = 1.33$, the Cherenkov threshold momentum and energy
 787 for various particles are given in Table 3.3. In contrast, γ -rays are detected
 788 indirectly via the combination of photons generated by Compton scattering
 789 and pair production. The threshold for detection in the SK detector is typically
 790 higher than the threshold for photon production. This is due to the fact that the
 791 attenuation of photons in the water means that typically $\sim 75\%$ of Cherenkov
 792 photons reach the ID PMTs. Then the collection and quantum efficiencies
 793 described in subsection 3.1.1 result in the number of detected photons being
 794 lower than the number of photons which reach the PMTs.

Particle	Threshold Momentum (MeV)	Threshold Energy (MeV)
Electron	0.5828	0.7751
Muon	120.5	160.3
Pion	159.2	211.7
Proton	1070.0	1423.1

Table 3.3: The threshold momentum and total energy for a particle to generate Cherenkov light in ultrapure water, as calculated in Equation 3.2 in ultrapure water which has refractive index $n = 1.33$.

795 The Frank-Tamm equation [106] describes the relationship between the num-
 796 ber of Cherenkov photons generated per unit length, dN/dx , the wavelength of
 797 the photons generated, λ , and the relativistic velocity of the charged particle,

$$\frac{d^2N}{dxd\lambda} = 2\pi\alpha \left(1 - \frac{1}{n^2\beta^2}\right) \frac{1}{\lambda^2}. \quad (3.4)$$

798 where α is the fine structure constant. For a 100MeV momentum electron,
 799 approximately 330 photons will be produced per centimeter in the $300\text{nm} \leq \lambda \leq$
 800 700nm region which the ID PMTs are most sensitive to [85].

801 3.2 The Tokai to Kamioka Experiment

802 The Tokai to Kamioka (T2K) experiment is a long-baseline neutrino oscillation
 803 experiment located in Japan. Proposed in the early 2000s [107, 108] to replace
 804 K2K [109], T2K was designed to observe electron neutrino appearance whilst
 805 precisely measuring the oscillation parameters associated with muon neutrino
 806 disappearance [110]. The experiment consists of a neutrino beam generated
 807 at the Japan Proton Accelerator Research Complex (J-PARC), a suite of near
 808 detectors situated 280m from the beam target, and the Super Kamiokande far
 809 detector positioned at a 295km baseline. The cross-section view of the T2K
 810 experiment is drawn in Figure 3.4.

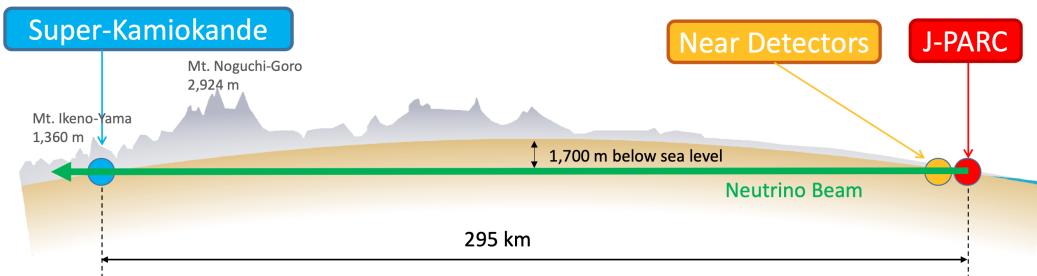


Figure 3.4: The cross-section view of the Tokai to Kamioka experiment illustrating the beam generation facility at J-PARC, the near detector situated at a baseline of 280m and the Super Kamiokande far detector situated 295km from the beam target.

811 The T2K collaboration makes world-leading measurements of the $\sin^2(\theta_{23})$,
 812 Δm_{32}^2 , and δ_{CP} oscillation parameters. Improvements in the precision and accu-
 813 racy of parameter estimates are still being made by including new data samples
 814 and developing the models which describe the neutrino interactions and detector
 815 responses [111]. Electron neutrino appearance was first observed at T2K in 2014
 816 [112] with 7.3σ significance.

817 The near detectors provide constraints on the beam flux and cross-section
 818 model parameters used within the oscillation analysis by observing the unoscil-
 819 lated neutrino beam. There are a host of detectors situated in the near detector hall
 820 (As illustrated in Figure 3.5): ND280 (subsection 3.2.3), INGRID (subsection 3.2.4),
 821 NINJA [113], WAGASCI [114], and Baby-MIND [115]. The latter three are not
 822 currently used within the oscillation analysis presented in this thesis.

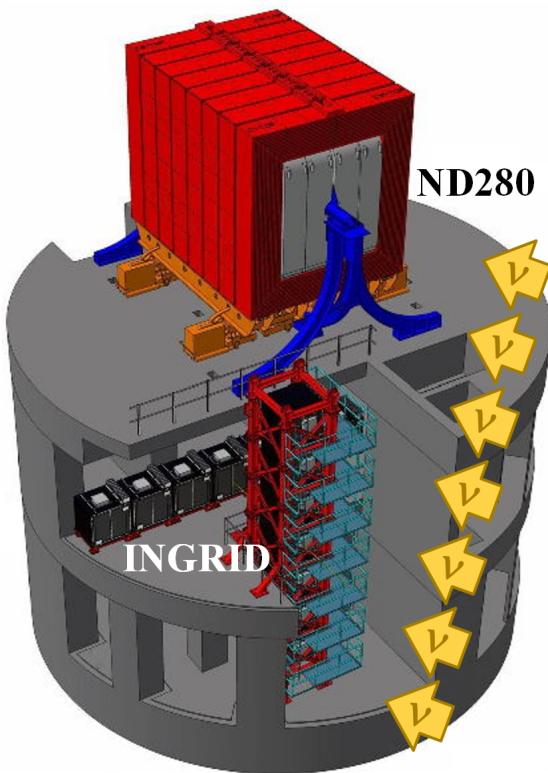


Figure 3.5: The near detector suite for the T2K experiment showing the ND280 and INGRID detectors. The distance between the detectors and the beam target is 280m.

823 Whilst this thesis presents the ND280 in terms of its purpose for the oscillation
 824 analysis, the detector can also make many cross-section measurements at neutrino
 825 energies of $O(1)$ GeV for the different targets within the detector [116, 117]. These
 826 measurements are of equal importance as they can lead the way in determining
 827 the model parameters used in the interaction models for the future high-precision
 828 era of neutrino physics.

3.2.1 Analysis Overview

There are two independent fitters, MaCh3 and BANFF, which perform the near detector fit. MaCh3 uses a bayesian Markov Chain Monte Carlo fitting technique, whereas BANFF uses a frequentist gradient descent technique. The output of each fitter is compared as a method of cross-checking the behaviour of the two fitters. This is done by comparing: the Monte Carlo predictions using various tunes, the likelihood that is calculated in each fitter and the post-fit constraint associated with every parameter used in the fit. Once validated, the output converted into a covariance matrix to describe the error and correlations between all the flux and cross-section parameters. This is then propagated to the far-detector oscillation analysis group.

The far detector group has three independent fitters: P-Theta, VALOR and MaCh3. The first two fitters use a hybrid frequentist fitting technique where the likelihood is minimised with respect to the parameters of interest and marginalised over all other parameters. These fitters use the covariance provided by the near detector fitters as a basis for implementing the near detector constraints. The MaCh3 fitter uses a simultaneous fit of all near and far detector samples. This removes any Gaussian assumptions when making the covariance matrix from the near detector results. The results for all three fitters are compares using a technique similar to the validation of the near detector fitters.

There are three particular tunes of the T2K flux and low energy cross section model typically considered. Firstly, the “generated” tune which is the set of dial values with which the Monte Carlo was generated. Secondly, the set of dial values which are taken from external data measurements and used as inputs. These are the “pre-fit” dial values. The reason these two sets of dial values are different is that the external data measurements are continually updated but it is very computationally intensive to regenerate a Monte Carlo prediction after each update. The final tune is the “post-fit”, “post-ND fit” or “post-BANFF” dial values. These are the values taken from the constraints provided by the near detector.

3.2.2 The Neutrino Beam

The neutrino beam used within the T2K experiment is described in [39, 43] and summarised below. The accelerator facility at J-PARC is composed of two sections; the primary and secondary beamlines. Figure 3.6 illustrates a schematic of the beamline, focusing mostly on the components of the secondary beamline. The primary beamline has three accelerators that progressively accelerate protons; a linear accelerator, a rapid-cycling synchrotron, and the main-ring (MR) synchrotron. Once fully accelerated by the MR, the protons have a kinetic energy of 30GeV. Eight bunches of these protons, separated by 500ns, are extracted per “spill” from the MR and directed towards a graphite target (a rod of length 91.4cm and diameter 2.6cm). Spills are extracted at 0.5Hz with $\sim 3 \times 10^{14}$ protons contained per spill.

The secondary beamline consists of three main components: the target station, the decay volume, and the beam dump. The target station is comprised of the target, beam monitors, and three magnetic focusing horns. The proton beam interacts with the graphite target to form a secondary beam of mostly pions and kaons. The secondary beam travels through a 96m long decay volume, generating neutrinos through the following decays [39],

$$\begin{array}{ll}
\pi^+ \rightarrow \mu^+ + \nu_\mu & \pi^- \rightarrow \mu^- + \bar{\nu}_\mu \\
K^+ \rightarrow \mu^+ + \nu_\mu & K^- \rightarrow \mu^- + \bar{\nu}_\mu \\
\rightarrow \pi^0 + e^+ + \nu_e & \rightarrow \pi^0 + e^- + \bar{\nu}_e \\
\rightarrow \pi^0 + \mu^+ + \nu_\mu & \rightarrow \pi^0 + \mu^- + \bar{\nu}_\mu \\
K_L^0 \rightarrow \pi^- + e^+ + \nu_e & K_L^0 \rightarrow \pi^+ + e^- + \bar{\nu}_e \\
\rightarrow \pi^- + \mu^+ + \nu_\mu & \rightarrow \pi^+ + \mu^- + \bar{\nu}_\mu \\
\mu^+ \rightarrow e^+ + \bar{\nu}_\mu + \nu_e & \mu^- \rightarrow e^- + \nu_\mu + \bar{\nu}_e.
\end{array}$$

The electrically charged component of the secondary beam is focused towards the far detector by the three magnetic horns. These horns direct charged particles of a particular polarity towards SK whilst defocusing the oppositely charged particles. This allows a mostly neutrino or mostly antineutrino beam to be used within the experiment, denoted as “forward horn current (FHC)” or “reverse horn current (RHC)” respectively.

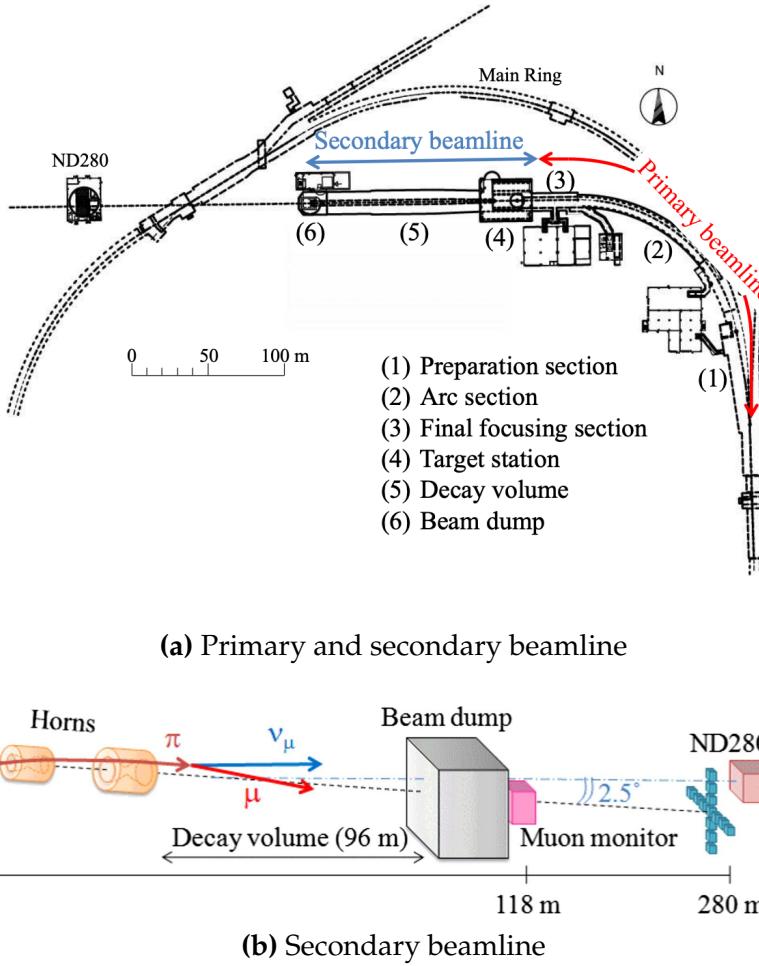


Figure 3.6: Top panel: Bird's eye view of the most relevant part of primary and secondary beamline used within the T2K experiment. The primary beamline is the main-ring proton synchrotron, kicker magnet, and graphite target. The secondary beamline consists of the three focusing horns, decay volume, and beam dump. Figure taken from [43]. Bottom panel: The side-view of the secondary beamline including the focusing horns, beam dump and neutrino detectors. Figure taken from [118].

884 Figure 3.7 illustrates the different contributions to the FHC and RHC neutrino
 885 flux. The low energy flux is dominated by the decay of pions whereas kaon
 886 decay becomes the dominant source of neutrinos for $E_\nu > 3\text{GeV}$. The “wrong-
 887 sign” component, which is the $\bar{\nu}_\mu$ background in a ν_μ beam, and the intrinsic
 888 irreducible ν_e background, are predominantly due to muon decay for $E_\nu <$
 889 2GeV . As the antineutrino production cross-section is smaller than the neutrino
 890 cross-section, the wrong-sign component is more dominant in the RHC beam
 891 as compared to that in the FHC beam.



Figure 3.7: The Monte Carlo prediction of the energy spectrum for each flavour of neutrino (ν_e , $\bar{\nu}_e$, ν_μ and $\bar{\nu}_\mu$) in the neutrino dominated beam FHC mode (Left) and antineutrino dominated beam RHC mode (Right) expected at SK. Taken from [119].

892 The beam dump, situated at the end of the decay volume, stops all charged
 893 particles other than highly energetic muons ($p_\mu > 5\text{GeV}$). The MuMon detector
 894 monitors the penetrating muons to determine the beam direction and inten-
 895 sity which is used to constrain some of the beam flux systematics within the
 896 analysis [118, 120].

897 The T2K experiment uses an off-axis beam to narrow the neutrino energy
 898 distribution. This was the first implementation of this technique in a long-
 899 baseline neutrino oscillation experiment after its original proposal [121]. Pion
 900 decay, $\pi \rightarrow \mu + \nu_\mu$, is a two-body decay. Consequently, the neutrino energy,
 901 E_ν , can be determined based on the pion energy, E_π , and the angle at which
 902 the neutrino is emitted, θ ,

$$E_\nu = \frac{m_\pi^2 - m_\mu^2}{2(E_\pi - p_\pi \cos(\theta))}, \quad (3.5)$$

903 where m_π and m_μ are the mass of the pion and muon respectively. For a fixed
 904 energy pion, the neutrino energy distribution is dependent upon the angle at
 905 which the neutrinos are observed from the initial pion beam direction. For the
 906 295km baseline at T2K, $E_\nu = 0.6\text{GeV}$ maximises the electron neutrino appearance
 907 probability, $P(\nu_\mu \rightarrow \nu_e)$, whilst minimising the muon disappearance probability,

908 $P(\nu_\mu \rightarrow \nu_\mu)$. Figure 3.8 illustrates the neutrino energy distribution for a range of
 909 off-axis angles, as well as the oscillation probabilities most relevant to T2K.

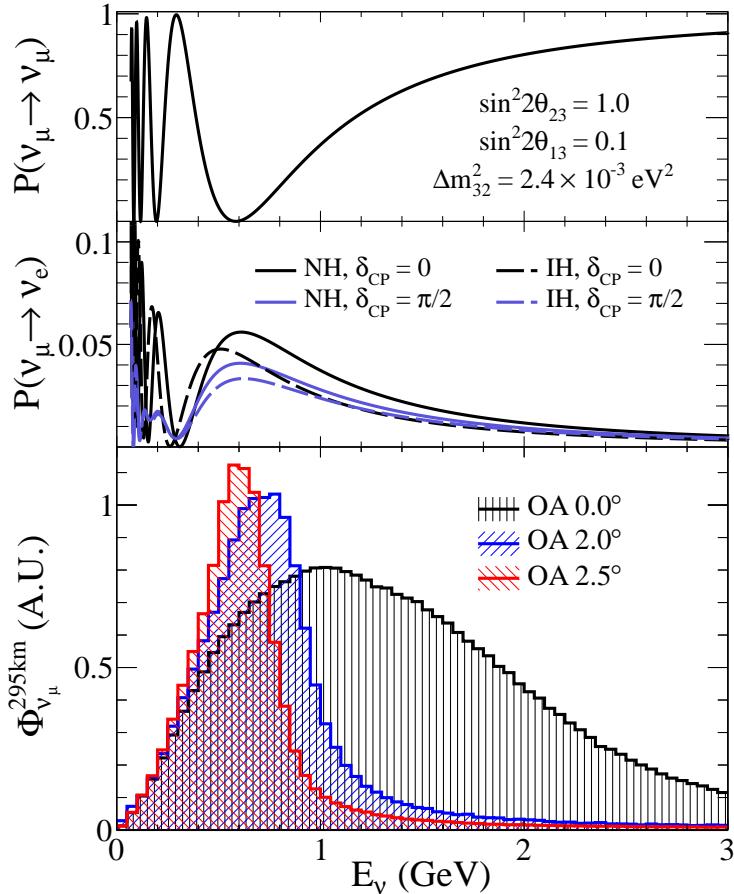


Figure 3.8: Top panel: T2K muon neutrino disappearance probability as a function of neutrino energy. Middle panel: T2K electron neutrino appearance probability as a function of neutrino energy. Bottom panel: The neutrino flux distribution for three different off-axis angles (Arbitrary units) as a function of neutrino energy.

910 3.2.3 The Near Detector at 280m

911 Whilst all the near detectors are situated in the same “pit” located at 280m from
 912 the beamline, the “ND280” detector is the off-axis detector which is situated at
 913 the same off-axis angle as the Super-Kamiokande far detector. It has two primary
 914 functions; firstly it measures the neutrino flux and secondly, it counts the event
 915 rates of different types of neutrino interactions. Both of these constrain the flux
 916 and cross-section systematics invoked within the model for a more accurate
 917 prediction of the expected event rate at the far detector.



Figure 3.9: The components of the ND280 detector. The neutrino beam travels from left to right. Taken from [43].

As illustrated in Figure 3.9, the ND280 detector consists of several sub-detectors. The most important part of the detector for this analysis is the tracker region. This is comprised of two-time projection chambers (TPCs) sandwiched between three fine grain detectors (FGDs). The FGDs contain both hydrocarbon plastics and water targets for neutrino interactions and provide track reconstruction near the interaction vertex. The emitted charged particles can then propagate into the TPCs which provide particle identification and momentum reconstruction. The FGDs and TPCs are further described in subsubsection 3.2.3.1 and subsubsection 3.2.3.2 respectively. The electromagnetic calorimeter (ECAL) encapsulates the tracker region alongside the π^0 detector (P0D). The ECAL measures the deposited energy from photons emitted from interactions within the FGD. The P0D constrains the cross-section of neutral current interactions which generate neutral pions, which is one of the largest backgrounds in the electron neutrino appearance oscillation channel. The P0D and ECAL detectors are detailed in subsubsection 3.2.3.3 and subsubsection 3.2.3.4 respectively. The entire detector is located within a large yoke magnet which produces a 0.2T magnetic field.

934 field. This design of the magnet also includes a scintillating detector called the
935 side muon range detector (SMRD), which is used to track high-angle muons as
936 well as acting as a cosmic veto. The SMRD is described in subsubsection 3.2.3.5.

937 **3.2.3.1 Fine Grained Detectors**

938 The T2K tracker region is comprised of two fine-grained detectors (FGD) and
939 three Time Projection Chambers (TPC). A detailed description of the FGD design,
940 construction, and assembly is found in [122] and summarised below. The FGDS
941 are the primary target for neutrino interactions with a mass of 1.1 tonnes per FGD.
942 Alongside this, the FGDS are designed to be able to track short-range particles
943 which do not exit the FGD. Typically, short-range particles are low momentum
944 and are observed as tracks that deposit a large amount of energy per unit length.
945 This means the FGD needs good granularity to resolve these particles. The
946 FGDS have the best timing resolution ($\sim 3\text{ns}$) of any of the sub-detectors of the
947 ND280 detector. As such, the FGDS are used for time of flight measurements
948 to distinguish forward-going positively charged particles from backward-going
949 negatively charged particles. Finally, any tracks which pass through multiple
950 sub-detectors are required to be track matched to the FGD.

951 Both FGDS are made from square scintillator planes of side length 186cm and
952 width 2.02cm. Each plane consists of two layers of 192 scintillator bars in an
953 X or Y orientation. A wavelength-shifting fiber is threaded through the center
954 of each bar and is read out by a multi-pixel photon counter (MPPC). FGD1 is
955 the most upstream of the two FGDS and contains 15 planes of carbon plastic
956 scintillator which is a common target in external neutrino scattering data. As
957 the far detector is a pure water target, 7 of the 15 scintillator planes in FGD2
958 have been replaced with a hybrid water-scintillator target. Due to the complexity
959 of the nucleus, nuclear effects can not be extrapolated between different nuclei.
960 Therefore having the ability to take data on one target which is the same as
961 external data and another target which is the same as the far detector target is
962 beneficial for reliable model parameter estimates.

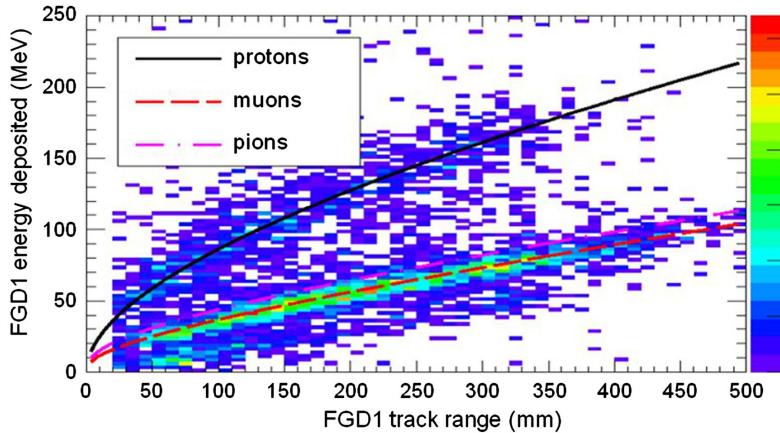


Figure 3.10: Comparison of data to Monte Carlo prediction of integrated deposited energy as a function of track length for particles that stopped in FGD1. Taken from [122].

963 The integrated deposited energy is used for particle identification. The FGD
964 can distinguish protons from other charged particles by comparing the integrated
965 deposited energy from data to Monte Carlo prediction as seen in Figure 3.10.

966 3.2.3.2 Time Projection Chambers

967 The majority of particle identification and momentum measurements within
968 ND280 are provided by three Time Projection Chambers (TPCs) [123]. The
969 TPCs are located on either side of the FGDs. They are located inside of the
970 magnetic field meaning the momentum of a charged particle can be determined
971 from the bending of the track.

972 Each TPC module consists of two gas-tight boxes, as shown in Figure 3.11,
973 which are made of non-magnetic material. The outer box is filled with CO₂ which
974 acts as an electrical insulator between the inner box and the ground. The inner box
975 forms the field cage which produces a uniform electric drift field of $\sim 275\text{V/cm}$
976 and is filled with an argon gas mixture. Charged particles moving through this
977 gas mixture ionize the gas and the ionised charge is drifted towards micromegas
978 detectors which measure the ionization charge. The time and position information
979 in the readout allows a three-dimensional image of the neutrino interaction.

980 The particle identification of tracks that pass through the TPCs is performed
981 using dE/dx measurements. Figure 3.12 illustrates the data to Monte Carlo



Figure 3.11: Schematic design of a Time Projection Chamber detector. Taken from [123].

982 distributions of the energy lost by a charged particle passing through the TPC as
 983 a function of the reconstructed particle momentum. The resolution is $7.8 \pm 0.2\%$
 984 meaning that electrons and muons can be distinguished. This allows reliable
 985 measurements of the intrinsic ν_e component of the beam.

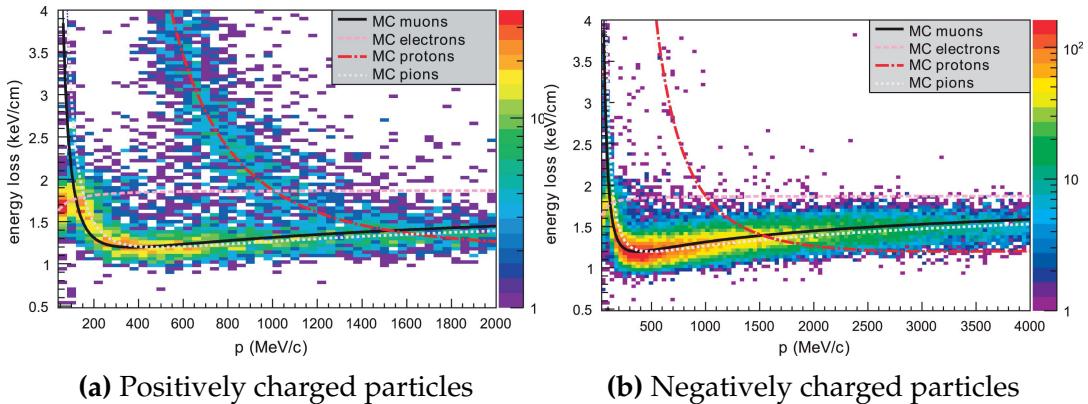


Figure 3.12: The distribution of energy loss as a function of reconstructed momentum for charged particles passing through the TPC, comparing data to Monte Carlo prediction. Taken from [123].

986 **3.2.3.3 π^0 Detector**

987 If one of the γ -rays from a $\pi^0 \rightarrow 2\gamma$ decay is missed at the far detector, the
 988 reconstruction will determine that event to be a charge current ν_e -like event.
 989 This is one of the main backgrounds hindering the electron neutrino appearance

990 searches. The π^0 detector (P0D) measures the cross-section of the neutral current
 991 induced neutral pion production on a water target to constrain this background.

992 The P0D is a cube of approximately 2.5m length consisting of layers of scin-
 993 tillating bars, brass and lead sheets, and water bags as illustrated in Figure 3.13.
 994 Two electromagnetic calorimeters are positioned at the most upstream and most
 995 downstream position in the sub-detector and the water target is situated in
 996 between them. The scintillator layers are built from two triangular bars orientated
 997 in opposite directions to form a rectangular layer. Each triangular scintillator bar
 998 is threaded with optical fiber which is read out by MPPCs. The high-Z brass and
 999 lead regions produce electron showers from the photons emitted in π^0 decay.

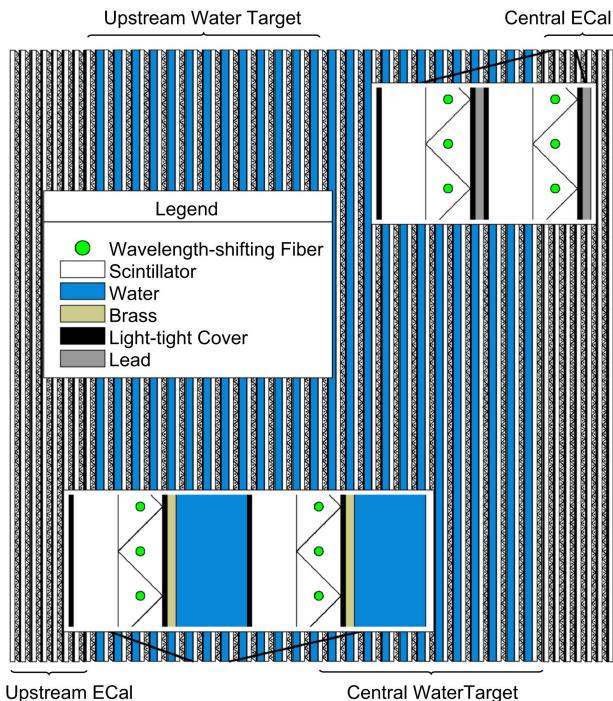


Figure 3.13: A schematic of the P0D side-view. Taken from [124].

1000 The sub-detector can generate measurements of NC1 π^0 cross-sections on a
 1001 water target by measuring the event rate both with and without the water target,
 1002 with the cross-section on a water target being determined as the difference. The to-
 1003 tal active mass is 16.1 tonnes when filled with water and 13.3 tonnes when empty.

1004 3.2.3.4 Electromagnetic Calorimeter

1005 The electromagnetic calorimeter [125] (ECal) encapsulates the P0D and tracking
 1006 sub-detectors. Its primary purpose is to aid π^0 reconstruction from any interac-
 1007 tion in the tracker. To do this, it measures the energy and direction of photon
 1008 showers from $\pi^0 \rightarrow 2\gamma$ decay. It can also distinguish pion and muon tracks
 1009 depending on the shape of the photon shower deposited.

1010 The ECal is comprised of three sections; the P0D ECal which surrounds the
 1011 P0D, the barrel ECal which encompasses the tracking region, and the downstream
 1012 ECal which is situated downstream of the tracker region. The barrel and down-
 1013 stream ECals are tracking calorimeters that focus on electromagnetic showers
 1014 from high-angle particles emitted from the tracking sub-detectors. Particularly in
 1015 the TPC, high-angle tracks (those which travel perpendicularly to the beam-axis)
 1016 can travel along a single scintillator bar resulting in very few hits. The width of
 1017 the barrel and downstream ECal corresponds to ~ 11 electron radiation lengths
 1018 to ensure a significant amount of the π^0 energy is contained. As the P0D has
 1019 its own calorimetry which reconstructs showers, the P0D ECal determines the
 1020 energy which escapes the P0D.

1021 Each ECal is constructed of multiple layers of scintillating bars sandwiched
 1022 between lead sheets. The scintillating bars are threaded with optical fiber and read
 1023 out by MPPCs. Each sequential layer of the scintillator is orientated perpendicular
 1024 to the previous which allows a three-dimensional event reconstruction. The
 1025 target mass of the P0D ECal, barrel ECal, and downstream ECal are 1.50, 4.80,
 1026 and 6.62 tonnes respectively.

1027 3.2.3.5 Side Muon Range Detector

1028 As illustrated in Figure 3.9, the ECal, FGDs, P0D, and TPCs are enclosed within
 1029 the UA1 magnet. Reconditioned after use in the UA1 [126] and NOMAD [127]
 1030 experiments, this magnet provides a uniform horizontal magnetic field of 0.2T
 1031 with an uncertainty of 2×10^{-4} T.

1032 Built into the UA1 magnet, the side muon range detector (SMRD)[128] monitors
1033 high-energy muons which leave the tracking region and permeate through
1034 the ECal. It additionally acts as a cosmic muon veto and trigger.

1035 **3.2.4 The Interactive Neutrino GRID**

1036 The Interactive Neutrino GRID (INGRID) detector is situated within the same
1037 “pit” as the other near detectors. It is aligned with the beam in the “on-axis”
1038 position and measures the beam direction, spread, and intensity. The detector
1039 was originally designed with 16 identical modules [43] (two modules have since
1040 been decommissioned) and a “proton” module. The design of the detector is 14
1041 modules oriented in a cross with length and height $10\text{m} \times 10\text{m}$, as illustrated
1042 in Figure 3.14.

1043 Each module is composed of iron sheets interlaced with eleven tracking
1044 scintillator planes for a total target mass of 7.1 tonnes per module. The scintillator
1045 design is an X-Y pattern of 24 bars in both orientations, where each bar contains
1046 wave-length shifting fibers which are connected to multi-pixel photon counters
1047 (MPPCs). Each module is encapsulated inside veto planes to aid the rejection
1048 of charged particles entering the module.

1049 The proton module is different from the other modules in that it consists
1050 of entirely scintillator planes with no iron target. The scintillator bars are also
1051 smaller than those used in the other modules to increase the granularity of
1052 the detector and improve tracking capabilities. The module sits in the center
1053 of the beamline and is designed to give precise measurements of quasi-elastic
1054 charged current interactions to evaluate the performance of the Monte Carlo
1055 simulation of the beamline.

1056 The INGRID detector can measure the beam direction to an uncertainty of
1057 0.4mrad and the beam centre within a resolution of 10cm [43]. The beam direction
1058 in both the vertical and horizontal directions is discussed in [129] and it is found
1059 to be in good agreement with the MUMON monitor described in subsection 3.2.2.



Figure 3.14: Left panel: The Interactive Neutrino GRID on-axis Detector. 14 modules are arranged in a cross-shape configuration, with the center modules being directly aligned with the on-axis beam. Right panel: The layout of a single module of the INGRID detector. Both figures are recreated from [43].

4

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Bayesian Statistics and Markov Chain Monte Carlo Techniques

1063 This thesis presents a Bayesian oscillation analysis. To extract the oscillation
1064 parameters, a Markov Chain Monte Carlo (MCMC) method is used. This chapter
1065 explains the theory of how parameter estimates can be determined using this
1066 technique and condenses the material found in the literature [130–133].

1067 The oscillation parameter determination presented here is built upon a si-
1068 multaneous fit to neutrino beam data in the near detector, beam data at SK, and
1069 atmospheric data at SK. In total, there are four oscillation parameters of interest
1070 ($\sin^2(\theta_{23})$, $\sin^2(\theta_{13})$, Δm_{32}^2 , and δ_{CP}), two oscillation parameters to which this
1071 study will not be sensitive ($\sin^2(\theta_{12})$, Δm_{21}^2) and many nuisance parameters that
1072 control the systematic uncertainty models.

1073 This analysis uses a Monte Carlo technique to generate a multi-dimensional
1074 probability distribution across all of the model parameters used in the fit. To
1075 determine an estimate for each parameter, this multi-dimensional object is in-
1076 tegrated over all other parameters. This process is called Marginalisation and
1077 is described in subsection 4.3.1. Monte Carlo techniques approximate the prob-
1078 ability distribution of each parameter within the limit of generating infinite
1079 samples. As ever, generating a large number of samples is time and resource-

1080 dependent. Therefore, an MCMC technique is utilised within this analysis to
1081 reduce the required number of steps to sufficiently sample the parameter space.
1082 This technique is described in further detail in subsection 4.2.1.

1083 The Bayesian analysis techniques used within this thesis are built within the
1084 MaCh3 framework [134]. This uses a custom MCMC library package exclusively
1085 supported and developed by the MaCh3 collaborators (which includes the author
1086 of this thesis).

1087 4.1 Bayesian Statistics

1088 Bayesian inference treats observable data, D , and model parameters, $\vec{\theta}$, on equal
1089 footing such that a probability model of both data and parameters is required.
1090 This is the joint probability distribution $P(D, \vec{\theta})$ and can be described by the
1091 prior distribution for model parameters $P(\vec{\theta})$ and the likelihood of the data given
1092 the model parameters $P(D|\vec{\theta})$,

$$P(D, \vec{\theta}) = P(D|\vec{\theta})P(\vec{\theta}). \quad (4.1)$$

1093 The prior distribution, $P(\vec{\theta})$, describes all previous knowledge about the
1094 parameters within the model. For example, if the risk of developing health
1095 problems is known to increase with age, the prior distribution would describe the
1096 increase. For the purpose of this analysis, the prior distribution is typically
1097 the best-fit values taken from external data measurements with a Gaussian
1098 uncertainty. The prior distribution can also contain correlations between model
1099 parameters. In an analysis using Monte Carlo techniques, the likelihood of
1100 measuring some data assuming some set of model parameters is calculated
1101 by comparing the Monte Carlo prediction generated at that particular set of
1102 model parameters to the data.

1103 It is parameter estimation that is important for this analysis and as such, Bayes'
1104 theorem [135] is applied to calculate the probability for each parameter to have a

1105 certain value given the observed data, $P(\vec{\theta}|D)$, which is known as the posterior
1106 distribution (often termed the posterior). This can be expressed as

$$P(\vec{\theta}|D) = \frac{P(D|\vec{\theta})P(\vec{\theta})}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}}. \quad (4.2)$$

1107 The denominator in Equation 4.2 is the integral of the joint probability distri-
1108 bution over all values of all parameters used within the fit. For brevity, the
1109 posterior distribution is

$$P(\vec{\theta}|D) \propto P(D|\vec{\theta})P(\vec{\theta}). \quad (4.3)$$

1110 For the purposes of this analysis, it is acceptable to neglect the normalisation
1111 term and focus on this proportional relationship.

1112 4.1.1 Application of Prior Knowledge

1113 The posterior distribution is proportional to the prior uncertainty applied to
1114 each parameter, as illustrated by Equation 4.3. This means that it is possible
1115 to change the prior after the posterior distribution has been determined. The
1116 prior uncertainty of a particular parameter can be ‘divided’ out of the posterior
1117 distribution and the resulting distribution can be reweighted using the new
1118 prior uncertainty that is to be applied. The methodology and implementation
1119 of changing the prior follows that described in [136].

1120 An example implementation that is useful for this analysis is the application
1121 of the “reactor constraint”. As discussed in section 2.4, an external constraint
1122 on $\sin^2(\theta_{13})$ is determined from measurements taken from reactor experiments.
1123 However, the sensitivities from just using the T2K and SK samples is equally
1124 as important. Without this technique, two fits would have to be run, doubling
1125 the required resources. Therefore, the key benefit for this analysis is the fact that
1126 only a single ‘fit’ has to be performed and can be used to build the two posterior
1127 distributions of the with and without reactor constraint applied.

1128 4.2 Monte Carlo Simulation

1129 Monte Carlo techniques are used to numerically solve a complex problem that
1130 does not necessarily have an analytical solution. These techniques rely on
1131 building a large ensemble of samples from an unknown distribution and then
1132 using the ensemble to approximate the properties of the distribution.

1133 An example that uses Monte Carlo techniques is to calculate the area under-
1134 neath a curve. For example, take the problem of calculating the area under a
1135 straight line with gradient $M = 0.4$ and intercept $C = 1.0$. Analytically, one can
1136 calculate the area under the line is equal to 30 units for $0 \leq x \leq 10$. Using Monte
1137 Carlo techniques, one can calculate the area under this line by throwing many
1138 random values for the x and y components of each sample and then calculating
1139 whether that point falls below the line. The area can then be calculated by the
1140 ratio of points below the line to the total number of samples thrown multiplied by
1141 the total area in which samples were scattered. The study is shown in Figure 4.1
1142 highlights this technique and finds the area under the curve to be 29.9 compared
1143 to an analytical solution of 30.0. The deviation of the numerical to analytical
1144 solution can be attributed to the number of samples used in the study. The
1145 accuracy of the approximation in which the properties of the Monte Carlo samples
1146 replicate those of the desired distribution is dependent on the number of samples
1147 used. Replicating this study with a differing number of Monte Carlo samples
1148 used in each study (As shown in Figure 4.2) highlights how the Monte Carlo
1149 techniques are only accurate within the limit of a high number of samples.

1150 Whilst the above example has an analytical solution, these techniques are just
1151 as applicable to complex solutions. Clearly, any numerical solution is only as
1152 useful as its efficiency. As discussed, the accuracy of the Monte Carlo technique is
1153 dependent upon the number of samples generated to approximate the properties
1154 of the distribution. Furthermore, if the positions at which the samples are
1155 evaluated are not ‘cleverly’ picked, the efficiency of the Monte Carlo technique
1156 significantly drops. Given the example in Figure 4.1, if the region in which the

1157 samples are scattered significantly extends passed the region of interest, many
1158 calculations will be calculated but do not add to the ability of the Monte Carlo
1159 technique to achieve the correct result. For instance, any sample evaluated at
1160 a $y \geq 5$ could be removed without affecting the final result. This does bring in
1161 an aspect of the ‘chicken and egg’ problem in that to achieve efficient sampling,
1162 one needs to know the distribution beforehand.

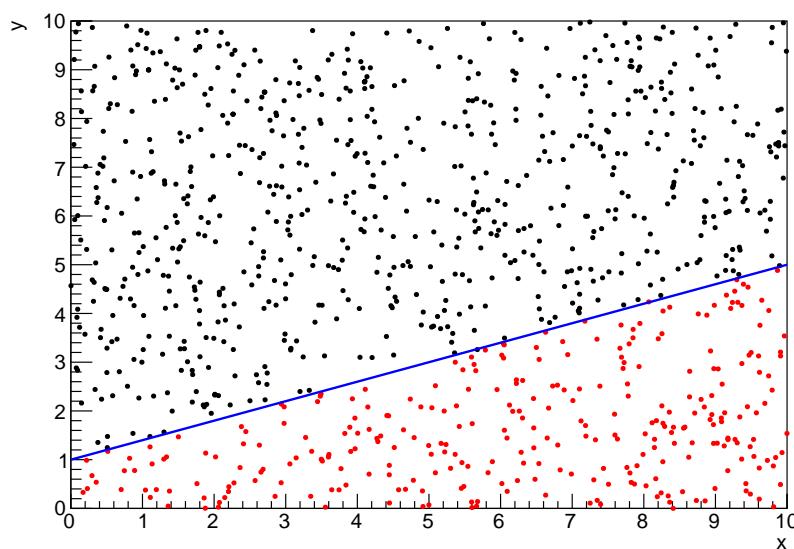


Figure 4.1: Example of using Monte Carlo techniques to find the area under the blue line. The gradient and intercept of the line are 0.4 and 1.0 respectively. The area found to be under the curve using one thousand samples is 29.9 units.

1163 4.2.1 Markov Chain Monte Carlo

1164 This analysis utilises a multi-dimensional probability distribution, with some
1165 dimensions being significantly more constrained than others. These constraints
1166 can be from prior knowledge of parameter distributions from external data or
1167 un-physical regions in which parameters can not exist. To maximise the efficiency
1168 of building the posterior distribution, a Markov Chain Monte Carlo (MCMC)
1169 technique is used. This employs a Markov chain to select the points at which
1170 to sample the posterior distribution. It performs a semi-random stochastic walk
1171 through the allowable parameter space. This builds a posterior distribution

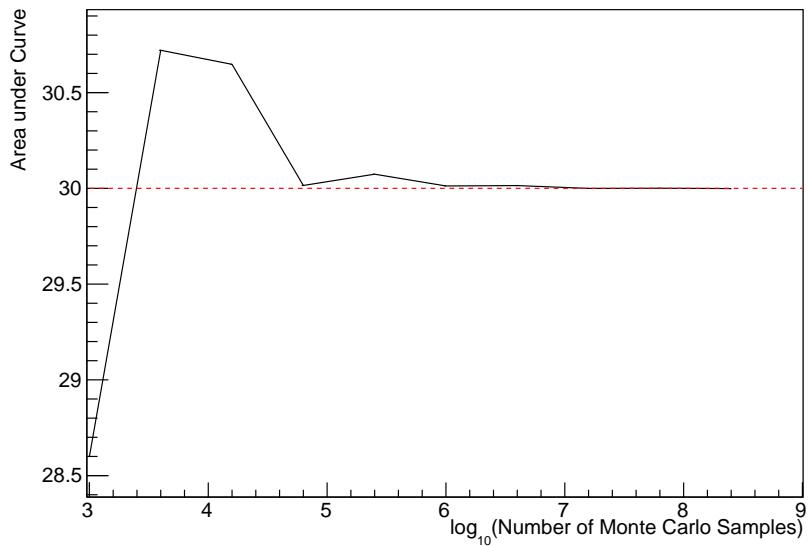


Figure 4.2: The area under a line of gradient 0.4 and intercept 1.0 for the range $0 \leq x \leq 10$ as calculated using Monte Carlo techniques as a function of the number of samples used in each repetition. The analytical solution to the area is 30 units as given by the red line.

1172 which has the property that the density of sampled points is proportional to the
 1173 probability density of that parameter. This means that the samples produced by
 1174 this technique are not statistically independent but they will cover the space
 1175 of the distribution.

1176 A Markov chain functions by selecting the position of step \vec{x}_{i+1} based on the
 1177 position of \vec{x}_i . The space in which the Markov chain selects samples is dependent
 1178 upon the total number of parameters utilised within the fit, where a discrete point
 1179 in this space is described by the N-dimensional space \vec{x} . In a perfectly operating
 1180 Markov chain, the position of the next step depends solely on the previous step
 1181 and not on the further history of the chain (\vec{x}_0, \vec{x}_1 , etc.). However, in solving
 1182 the multi-dimensionality of the fit used within this analysis, each step becomes
 1183 correlated with several of the steps preceding itself. Providing the MCMC chain is
 1184 well optimised, it will begin to converge towards a unique stationary distribution.
 1185 The period between the chain's initial starting point and the convergence to the
 1186 unique stationary distribution is colloquially known as the burn-in period. Once
 1187 the chain reaches the stationary distribution, all points sampled after that point

₁₁₈₈ will look like samples from that distribution.

₁₁₈₉ Further details of the theories underpinning MCMC techniques are discussed
₁₁₉₀ in [131] but can be summarised by the requirement that the chain satisfies the
₁₁₉₁ three ‘regularity conditions’:

- ₁₁₉₂ • Irreducibility: From every position in the parameter space \vec{x} , there must
₁₁₉₃ exist a non-zero probability for every other position in the parameter space
₁₁₉₄ to be reached.
- ₁₁₉₅ • Recurrence: Once the chain arrives at the stationary distribution, every step
₁₁₉₆ following from that position must be samples from the same stationary
₁₁₉₇ distribution.
- ₁₁₉₈ • Aperiodicity: The chain must not repeat the same sequence of steps at any
₁₁₉₉ point throughout the sampling period.

₁₂₀₀ The output of the chain after burn-in (i.e. the sampled points after the chain
₁₂₀₁ has reached the stationary distribution) can be used to approximate the posterior
₁₂₀₂ distribution and model parameters $\vec{\theta}$. To achieve the requirement that the unique
₁₂₀₃ stationary distribution found by the chain be the posterior distribution, one
₁₂₀₄ can use the Metropolis-Hastings algorithm. This guides the stochastic process
₁₂₀₅ depending on the likelihood of the current proposed step compared to that
₁₂₀₆ of the previous step.

₁₂₀₇ 4.2.2 Metropolis-Hastings Algorithm

₁₂₀₈ As a requirement for MCMCs, the Markov chain implemented in this technique
₁₂₀₉ must have a unique stationary distribution that is equivalent to the posterior
₁₂₁₀ distribution. To ensure this requirement and that the regularity conditions are
₁₂₁₁ met, this analysis utilises the Metropolis-Hastings (MH) algorithm [137, 138].
₁₂₁₂ For the i^{th} step in the chain, the MH algorithm determines the position in the
₁₂₁₃ parameter space to which the chain moves to based on the current step, \vec{x}_i , and
₁₂₁₄ the proposed step, \vec{y}_{i+1} . The proposed step is randomly selected from some

1215 proposal function $f(\vec{x}_{i+1}|\vec{x}_i)$, which depends solely on the current step (ie. not
1216 the further history of the chain). The next step in the chain \vec{x}_{i+1} can be either the
1217 current step or the proposed step determined by whether the proposed step is
1218 accepted or rejected. To decide if the proposed step is selected, the acceptance
1219 probability, $\alpha(\vec{x}_i, \vec{y}_i)$, is calculated as

$$\alpha(\vec{x}_i, \vec{y}_{i+1}) = \min\left(1, \frac{P(\vec{y}_{i+1}|D)f(\vec{x}_i|\vec{y}_{i+1})}{P(\vec{x}_i|D)f(\vec{y}_{i+1}|\vec{x}_i)}\right). \quad (4.4)$$

1220 Where $P(\vec{y}_{i+1}|D)$ is the posterior distribution as introduced in section 4.1. To
1221 simplify this calculation, the proposal function is required to be symmetric such
1222 that $f(\vec{x}_i|\vec{y}_{i+1}) = f(\vec{y}_{i+1}|\vec{x}_i)$. In practice, a multi-variate Gaussian distribution
1223 centered on \vec{x}_i is used to throw parameter proposals. This reduces Equation 4.4 to

$$\alpha(\vec{x}_i, \vec{y}_{i+1}) = \min\left(1, \frac{P(\vec{y}_{i+1}|D)}{P(\vec{x}_i|D)}\right). \quad (4.5)$$

1224 After calculating this quantity, a random number, β , is generated uniformly
1225 between 0 and 1. If $\beta \leq \alpha(\vec{x}_i, \vec{y}_{i+1})$, the proposed step is accepted. Otherwise,
1226 the chain sets the next step equal to the current step. This procedure is repeated
1227 for subsequent steps. This can be interpreted as if the posterior probability
1228 of the proposed step is greater than that of the current step, ($P(\vec{y}_{i+1}|D) \geq$
1229 $P(\vec{x}_i|D)$), the proposed step will always be accepted. If the opposite is true,
1230 ($P(\vec{y}_{i+1}|D) \leq P(\vec{x}_i|D)$), the proposed step will be accepted with probability
1231 $P(\vec{x}_i|D)/P(\vec{y}_{i+1}|D)$. This ensures that the Markov chain does not get trapped
1232 in any local minima in the potentially non-Gaussian posterior distribution. The
1233 outcome of this technique is that the density of steps taken in a discrete region
1234 is directly proportional to the probability density in that region.

1235 4.2.3 MCMC Optimisation

1236 As discussed in subsection 4.2.2, the proposal function invoked within the MH
1237 algorithm can take any form and the chain will still converge to the stationary
1238 distribution. At each set of proposed parameter values, a prediction of the same
1239 spectra has to be generated which requires significant computational resources.

1240 Therefore, the number of steps taken before the unique stationary distribution
1241 is found should be minimised as only steps after convergence add information
1242 to the oscillation analysis. Furthermore, the chain should entirely cover the
1243 allowable parameter space to ensure that all values have been considered. Tuning
1244 the distance that the proposal function jumps between steps on a parameter-by-
1245 parameter basis can both minimise the length of the burn-in period and ensure
1246 that the correlation between step \vec{x}_i and \vec{x}_j is sufficiently small.

1247 The effect of changing the width of the proposal function is highlighted in
1248 Figure 4.3. Three scenarios, each with the same underlying stationary distribution
1249 (A Gaussian of width 1.0 and mean 0.), are presented. The only difference between
1250 the three scenarios is the width of the proposal function, colloquially known as
1251 the ‘step size σ ’. Each scenario starts at an initial parameter value of 10.0 which
1252 would be considered an extreme variation. For the case where $\sigma = 0.1$, it is
1253 clear to see that the chain takes a long time to reach the expected region of the
1254 parameter. This indicates that this chain would have a large burn-in period and
1255 does not converge to the stationary distribution until step ~ 500 . Furthermore,
1256 whilst the chain does move towards the expected region, each step is significantly
1257 correlated with the previous. Considering the case where $\sigma = 5.0$, the chain
1258 approaches the expected parameter region almost instantly meaning that the
1259 burn-in period is not significant. However, there are clearly large regions of steps
1260 where the chain does not move. This is likely due to the chain proposing steps
1261 in the tails of the distribution which have a low probability of being accepted.
1262 Consequently, this chain would take a significant number of steps to fully span
1263 the allowable parameter region. For the final scenario, where $\sigma = 0.5$, you can
1264 see a relatively small burn-in period of approximately 100 steps. Once the chain
1265 reaches the stationary distribution, it moves throughout the expected region of
1266 parameter values many times, sufficiently sampling the full parameter region.
1267 This example is a single parameter varying across a continuous distribution and
1268 does not fully reflect the difficulties in the many-hundred multi-variate parameter

1269 distribution used within this analysis. However, it does give a conceptual idea of
1270 the importance of selecting the proposal function and associated step size.



Figure 4.3: Three MCMC chains, each with a stationary distribution equal to a Gaussian centered at 0 and width 1 (As indicated by the black dotted lines). All of the chains use a Gaussian proposal function but have different widths (or ‘step size σ ’). The top panel has $\sigma = 0.1$, middle panel has $\sigma = 0.5$ and the bottom panel has $\sigma = 5.0$.

1271 As discussed, step size tuning directly correlates to the average step accep-
1272 tance rate. If the step size is too small, many steps will be accepted but the
1273 chain moves slowly. If the opposite is true, many steps will be rejected as the
1274 chain proposes steps in the tails of the distribution. Discussion in [139] suggests
1275 that the ‘ideal’ acceptance rate of a high dimension MCMC chain should be

1276 approximately $\sim 25\%$. An “ideal” step size [139] of

$$\sigma = \frac{2.4}{N_p}, \quad (4.6)$$

1277 where N_p is the number of parameters included in the MCMC fit. However,
1278 the complex correlations between systematics mean that some parameters have
1279 to be hand-tuned and many efforts have been taken to select a set of parameter-
1280 by-parameter step sizes to approximately reach the ideal acceptance rate.

1281 Figure 4.4 highlights the likelihood as calculated by the fit in subsection 8.2.4
1282 as a function of the number of steps in each chain. In practice, many independent
1283 MCMC chains are run simultaneously to parallelise the task of performing the
1284 fit. This figure overlays the distribution found in each chain. As seen, the
1285 likelihood decreases from its initial value and converges towards a stationary
1286 distribution after $\sim 1 \times 10^5$ steps.

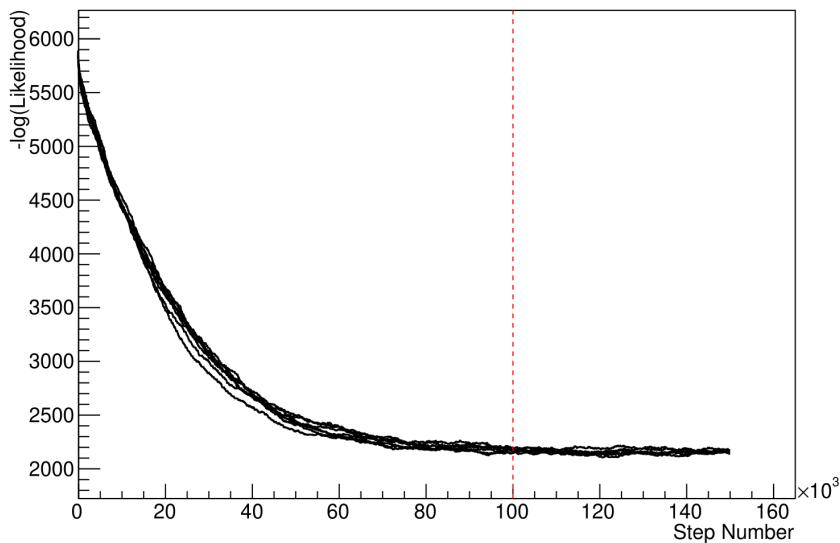


Figure 4.4: The log-likelihood from the fit detailed in subsection 8.2.4 as a function of the number of steps accumulated in each fit. Many independent MCMC chains were run in parallel and overlaid on this plot. The red line indicates the 1×10^5 step burn-in period after which the log-likelihood becomes stable.

1287 Multiple configurations of this analysis have been performed throughout this
1288 thesis where different samples or systematics have been used. For all of these con-
1289 figurations, it was found that a burnin period of 1×10^5 was sufficient in all cases.

1290 4.3 Understanding the MCMC Results

1291 The previous sections have described how to generate the posterior probability
1292 distribution using Bayesian MCMC techniques. However, this analysis focuses
1293 on oscillation parameter determination. The posterior distribution output from
1294 the chain is a high-dimension object, with as many dimensions as there are
1295 parameters included in the oscillation analysis. However, this multi-dimensional
1296 object is difficult to conceptualize so parameter estimations are often presented
1297 in one or two-dimensional projections of this probability distribution. To do
1298 this, marginalisation techniques are invoked.

1299 4.3.1 Marginalisation

1300 The output of the MCMC chain is a highly dimensional probability distribution
1301 which is very difficult to interpret. From the standpoint of an oscillation analysis
1302 experiment, the one or two-dimensional ‘projections’ of the oscillation parameters
1303 of interest are most relevant. Despite this, the best fit values and uncertainties on
1304 the oscillation parameters of interest should correctly encapsulate the correlations
1305 to the other systematic uncertainties (colloquially called ‘nuisance’ parameters).
1306 For this joint beam and atmospheric analysis, the oscillation parameters of
1307 interest are $\sin^2(\theta_{23})$, $\sin^2(\theta_{13})$, Δm_{32}^2 , and δ_{CP} . All other parameters (includ-
1308 ing the oscillation parameters this fit is insensitive to) are deemed nuisance
1309 parameters. To generate these projections, the posterior distribution is integrated
1310 over all nuisance parameters. This is called marginalisation. This technique
1311 also explains why it is acceptable to neglect the normalisation constant of the
1312 posterior distribution, which was discussed in section 4.1.

1313 A simple example of the marginalisation technique is to imagine the scenario
1314 where two coins are flipped. To determine the probability that the first coin
1315 returned a ‘head’, the exact result of the second coin flip is disregarded and
1316 simply integrated over. For the parameters of interest, $\vec{\theta}_i$, the marginalised

1317 posterior is calculated by integrating over the nuisance parameters, $\vec{\theta}_n$. In this
1318 case, Equation 4.2 becomes

$$P(\vec{\theta}_i|D) = \frac{\int P(D|\vec{\theta}_i, \vec{\theta}_n)P(\vec{\theta}_i, \vec{\theta}_n)d\vec{\theta}_n}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}}. \quad (4.7)$$

1319 Where $P(\vec{\theta}_i, \vec{\theta}_n)$ encodes the prior knowledge about the uncertainty and
1320 correlations between the parameters of interest and the nuisance parameters.
1321 In practice, this is simply taking the one or two-dimensional projection of the
1322 multi-dimensional probability distribution.

1323 While in principle an easy solution to a complex problem, correlations be-
1324 tween the interesting and nuisance parameters can bias the marginalised results.
1325 A similar effect is found when the parameters being marginalised over have
1326 non-Gaussian probability distributions. For example, Figure 4.5 highlights the
1327 marginalisation bias in the probability distribution found for a parameter when
1328 requiring a correlated parameter to have a positive parameter value. Due to
1329 the complex nature of the oscillation parameter fit presented in this thesis, there
1330 are correlations occurring between the oscillation parameters of interest and the
1331 other nuisance parameters included in the fit.



Figure 4.5: Left: The two-dimensional probability distribution for two correlated parameters x and y . The red distribution shows the two-dimensional probability distribution when $0 \leq x \leq 5$. Right: The marginalised probability distribution for the y parameter found when requiring the x to be bound between $-5 \leq x \leq 5$ and $0 \leq x \leq 5$ for the black and red distribution, respectively.

4.3.2 Parameter Estimation and Credible Intervals

The purpose of this analysis is to determine the best fit values for the oscillation parameters that the beam and atmospheric samples are sensitive to: $\sin^2(\theta_{23})$, $\sin^2(\theta_{13})$, Δm_{32}^2 , and δ_{CP} . The posterior probability density, taken from the output MCMC chain, is binned in these parameters. The parameter best-fit point is then taken to be the value that has the highest posterior probability. This is performed in both one and two-dimensional projections.

However, the single best-fit point in a given parameter is not of much use on its own. The uncertainty on the best-fit point must also be presented using credible intervals. The definition of the 1σ credible interval is that there is 68% belief that the parameter is within those bounds. For a more generalised definition, the credible interval is the region, R , of the posterior distribution that contains a specific fraction of the total probability, such that

$$\int_R P(\theta|D)d\theta = \alpha. \quad (4.8)$$

Where θ is the parameter being evaluated. This technique then calculates the $\alpha \times 100\%$ credible interval.

In practice, this analysis uses the highest posterior density (HPD) credible intervals which are calculated through the following method. First, the probability distribution is area-normalised such that it has an integrated area equal to 1.0. The bins of probability are then summed from the highest to lowest until the sum exceeds the 1σ level (0.68 in this example). This process is repeated for a range of credible intervals, notably the 1σ , 2σ and 3σ along with other levels where the critical values for each level can be found in [76]. This process can be repeated for the two-dimensional probability distributions by creating two-dimensional contours of credible intervals rather than a one-dimensional result.

1356 4.3.3 Bayesian Model Comparisons

1357 Due to the matter resonance, this analysis has some sensitivity to the mass
 1358 hierarchy of neutrino states (whether Δm_{32}^2 is positive or negative) and the
 1359 octant of $\sin^2(\theta_{23})$. The Bayesian approach utilised within this analysis gives an
 1360 intuitive method of model comparison by determining which hypothesis is most
 1361 favourable. Taking the ratio of Equation 4.3 for the two hypotheses of normal
 1362 hierarchy, NH , and inverted hierarchy, IH , gives

$$\frac{P(\vec{\theta}_{NH}|D)}{P(\vec{\theta}_{IH}|D)} = \frac{P(D|\vec{\theta}_{NH})}{P(D|\vec{\theta}_{IH})} \times \frac{P(\vec{\theta}_{NH})}{P(\vec{\theta}_{IH})}. \quad (4.9)$$

1363 The middle term defines the Bayes factor, $B(NH/IH)$, which is a data-driven
 1364 interpretation of how strong the data prefers one hierarchy to the other. For this
 1365 analysis, equal priors on both mass hierarchy hypotheses are chosen ($P(\vec{\theta}_{NH}) =$
 1366 $P(\vec{\theta}_{IH}) = 0.5$). In practice, the MCMC chain proposes a value of $|\Delta m_{32}^2|$ and
 1367 then applies a 50% probability that the value is sign flipped. Consequently,
 1368 the Bayes factor can be calculated from the ratio of the probability density in
 1369 either hypothesis. This equates to counting the number of steps taken in the
 1370 normal and inverted hierarchies and taking the ratio. The same approach can be
 1371 taken to compare the upper octant (UO) compared to the lower octant (LO)
 1372 hypothesis of $\sin^2(\theta_{23})$.

$\log_{10}(B_{AB})$	B_{AB}	Strength of Preference
< 0.0	< 1	No preference for hypothesis A (Supports hypothesis B)
0.0 – 0.5	1.0 – 3.16	Preference for hypothesis A is weak
0.5 – 1.0	3.16 – 10.0	Preference for hypothesis A is substantial
1.0 – 1.5	10.0 – 31.6	Preference for hypothesis A is strong
1.5 – 2.0	31.6 – 100.0	Preference for hypothesis A is very strong
> 2.0	> 100.0	Decisive preference for hypothesis A

Table 4.1: Jeffreys scale for strength of preference for two models A and B as a function of the calculated Bayes factor ($B_{AB} = B(A/B)$) between the two models [140]. The original scale is given in terms of $\log_{10}(B(A/B))$ but converted to linear scale for easy comparison throughout this thesis.

1373 Whilst the value of the Bayes factor should always be shown, the Jeffreys scale
 1374 [140] (highlighted in Table 4.1) gives an indication of the strength of preference

1375 for one model compared to the other. Other interpretations of the strength of
1376 preference of a model exist, e.g. the Kass and Raferty Scale [141].

1377 4.3.4 Comparison of MCMC Output to Expectation

1378 To ensure the fit is performing well, a best-fit spectrum is produced using the
1379 posterior probability distribution and compared with the data, allowing easy
1380 by-eye comparisons to be made. A simple method of doing this is to perform a
1381 comparison in the fitting parameters (e.g. the reconstructed neutrino energy for
1382 T2K far detector beam samples) of the spectra generated by the MCMC chain to
1383 ‘data’. This ‘data’ could be true data or some variation of Monte Carlo prediction.
1384 This allows easy comparison of the MCMC probability distribution to the data. To
1385 perform this, N steps from the post-burnin MCMC chain are randomly selected.
1386 From these, the Monte Carlo prediction at each step is generated by reweighting
1387 the model parameters to the values specified at that step. Due to the probability
1388 density being directly correlated with the density of steps in a certain region,
1389 parameter values close to the best fit value are most likely to be selected.

1390 In practice, for each bin of the fitting parameters has a probability distribution
1391 of event rates, with one entry per sampled MCMC step. This distribution is
1392 binned where the bin with the highest probability is selected as the mean and an
1393 error on the width of this probability distribution is calculated using the approach
1394 highlighted in subsection 4.3.2. Consequently, the best fit distribution in the fit
1395 parameter is not necessarily that which would be attained by reweighting the
1396 Monte Carlo prediction to the most probable parameter values.

1397 A similar study can be performed to illustrate the freedom of the model
1398 parameter space prior to the fit. This can be done by throwing parameter values
1399 from the prior uncertainty of each parameter.

5

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1402

Simulation, Reconstruction, and Event Reduction

1403 As a crucial part of the oscillation analysis, an accurate prediction of the expected
1404 neutrino spectrum at the far detector is required. This includes modeling the
1405 flux generation, neutrino interactions, and detector effects. All of the simulation
1406 packages required to do this are briefly described in section 5.1. The reconstruc-
1407 tion of neutrino events in the far detector, including the `fitQun` algorithm, is
1408 documented in section 5.2. This also includes data quality checks of the SK-V
1409 data which the author performed for the T2K oscillation analysis presented at the
1410 Neutrino 2020 conference [75]. Finally, section 5.3 describes the steps taken in the
1411 SK detector to trigger on events of interest whilst removing the comparatively
1412 large rate of cosmic ray muon events.

1413 5.1 Simulation

1414 In order to generate a Monte Carlo prediction of the expected event rate at
1415 the far detector, all the processes in the beam and atmospheric fluxes, neutrino
1416 interaction, and detector need to be modeled.

1417 5.1.1 Neutrino Flux

1418 The beamline simulation consists of three distinct parts: the initial hadron
1419 interaction modeled by FLUKA [142], the target station geometry and particle
1420 tracking performed by JNUBEAM, [143, 144] and any hadronic re-interactions
1421 simulated by GCALOR [145]. The primary hadronic interactions are $O(10)\text{GeV}$,
1422 where FLUKA matches external cross-section data better than GCALOR [146].
1423 However, FLUKA is not very adaptable so a small simulation is built to model
1424 the interactions in the target and the output is then passed to JNUBEAM and
1425 GCALOR for propagation. The hadronic interactions are tuned to data from the
1426 NA61/SHINE [147–149] and HARP [150] experiments. The tuning is done by
1427 reweighting the FLUKA and GCALOR predictions to match the external data
1428 multiplicity and cross-section measurements, based on final state particle kine-
1429 matics [146]. The culmination of this simulation package generates the predicted
1430 flux for neutrino and antineutrino beam modes which are illustrated in Figure 3.7.

1431 The atmospheric neutrino flux is simulated by the HKKM model [51, 53]. The
1432 primary cosmic ray flux is tuned to AMS [151] and BESS [152] data assuming
1433 the US-standard atmosphere '76 [153] density profile and includes geomagnetic
1434 field effects. The primary cosmic rays interact to generate pions and muons.
1435 The interaction of these secondary particles to generate neutrinos is handled by
1436 DPMJET-III [154] for energies above 32GeV and JAM [53, 155] for energies below
1437 that value [49]. These hadronic interactions are tuned to BESS and L3 data [156,
1438 157] using the same methodology as the tuning of the beamline simulation. The
1439 energy and cosine zenith predictions of $\nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$ flux are given in Figure 2.3
1440 and Figure 2.5, respectively. The flux is approximately symmetrical and peaked
1441 around the horizon ($\cos(\theta_Z) = 0.0$). This is because horizontally-going pions
1442 and kaons can travel further than their vertically-going counterparts resulting
1443 in a larger probability of decaying to neutrinos. The symmetry is broken in
1444 lower-energy neutrinos due to geomagnetic effects, which modify the track of the
1445 primary cosmic rays. Updates to the HKKM model are currently ongoing [158].

5.1.2 Neutrino Interaction

Once a flux prediction has been made for all three detectors, NEUT 5.4.0 [159, 160] models the interactions of the neutrinos in the detectors. For the purposes of this analysis, quasi-elastic (QE), meson exchange (MEC), single meson production (PROD), coherent pion production (COH), and deep inelastic scattering (DIS) interactions are simulated. These interaction categories can be further broken down by whether they were propagated via a W^\pm boson in Charged Current (CC) interactions or via a Z^0 boson in Neutral Current (NC) interactions. CC interactions have a charged lepton in the final state, which can be flavour-tagged in reconstruction to determine the flavour of the neutrino. In contrast, NC interactions have a neutrino in the final state so no flavour information can be determined from the observables left in the detector after an interaction. This is the reason why neutrinos that interact through NC modes are assumed to not oscillate within this analysis. Both CC and NC interactions are modeled for all the above interaction categories, other than MEC interactions which are only modeled for CC events.

As illustrated in Figure 5.1, CCQE interactions dominate the cross-section of neutrino interactions around $E_\nu \sim 0.5\text{GeV}$. The NEUT implementation adopts the Llewellyn Smith [161] model for neutrino-nucleus interactions, where the nuclear ground state of any bound nucleons (neutrino-oxygen interactions) is approximated by a spectral-function [162] model that simulates the effects of Fermi momentum and Pauli blocking. The cross-section of QE interactions is controlled by vector and axial-vector form factors parameterised by the BBBA05 [163] model and a dipole form factor with $M_A^{QE} = 1.21\text{GeV}$ fit to external data [164], respectively. NEUT implements the Valencia [165] model to simulate MEC events, where two nucleons and two holes in the nuclear target are produced (often called 2p2h interactions).

For neutrinos of energy $O(1)\text{GeV}$, PROD interactions become dominant. These predominantly produce charged and neutral pions although γ , kaon,



Figure 5.1: The NEUT prediction of the ν_μ -H₂O cross-section overlaid on the T2K ν_μ flux. The charged current (black, solid) and neutral current (black, dashed) inclusive, charged current quasi-elastic (blue, solid), charged current 2p2h (blue, dashed), charged current single pion production (pink), and charged current multi- π and DIS (Purple) cross-sections are illustrated. Figure taken from [159].

and η production is also considered. To simulate these interactions, the Berger-Sehgal [166] model is implemented within NEUT. It simulates the excitation of a nucleon from a neutrino interaction, production of an intermediate baryon, and the subsequent decay to a single meson or γ . Pions can also be produced through COH interactions, which occur when the incoming neutrino interacts with the entire oxygen nucleus leaving a single pion outside of the nucleus. NEUT utilises the Berger-Sehgal [167] model to simulate these COH interactions.

DIS and multi- π producing interactions become the most dominant for energies $> O(5)\text{GeV}$. PYTHIA [168] is used to simulate any interaction with invariant mass $W > 2\text{GeV}/c^2$, which produces at least one meson. For any interaction which produces at least two mesons but has $W < 2\text{GeV}/c^2$, the

¹⁴⁸⁶ Bronner model is used [169]. Both of these models use Parton distribution
¹⁴⁸⁷ functions based on the Bodek-Yang model [170–172].

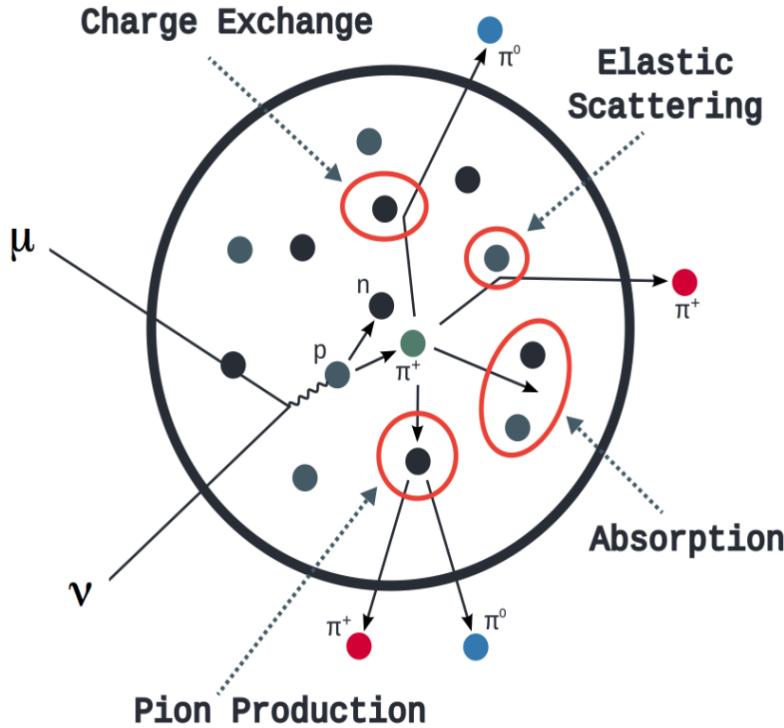


Figure 5.2: Illustration of the various processes which a pion can undergo before exiting the nucleus. Taken from [173].

¹⁴⁸⁸ Any pion that is produced within the nucleus can re-interact through final
¹⁴⁸⁹ state interactions before it exits, as illustrated by the scattering, absorption,
¹⁴⁹⁰ production, and exchange interactions in Figure 5.2. These re-interactions alter
¹⁴⁹¹ the observable particles within the detector. For instance, if the charged pion
¹⁴⁹² from a CC PROD interaction is absorbed, the observables would mimic a CC QE
¹⁴⁹³ interaction. To simulate these effects, NEUT uses a semi-classical intranuclear
¹⁴⁹⁴ cascade model [159]. This cascade functions by stepping the pion through the
¹⁴⁹⁵ nucleus in fixed-length steps equivalent to $dx = R_N/100$, where R_N is the radius
¹⁴⁹⁶ of the nucleus. At each step, the simulation allows the pion to interact through
¹⁴⁹⁷ scattering, charged exchange, absorption, or production with an interaction-
¹⁴⁹⁸ dependent probability calculated from a fit to external data [174]. This cascade
¹⁴⁹⁹ continues until the pion is absorbed or exits the nucleus.

1500 **5.1.3 Detector**

1501 Once the final state particle kinematics have been determined by NEUT, they
1502 are passed into the detector simulation. The near detectors, ND280 and INGRID,
1503 are simulated using a GEANT4 package [43, 175] to simulate the detector geom-
1504 etry, particle tracking, and energy deposition. The response of the detectors is
1505 simulated using the elecSim package [43].

1506 The far detector simulation is based upon the original Kamiokande experi-
1507 ment software which uses the GEANT3-based SKDETSIM [43, 176] package. This
1508 simulates the interactions of particles in the water as well as Cherenkov light
1509 production. The water quality and PMT calibration measurements detailed in
1510 subsection 3.1.2 are also used within this simulation to make accurate predictions
1511 of the detector response.

1512 Any event which generates optical photons that occurs in SK will be observed
1513 by the PMT array, where each PMT records the time and accumulated charge.
1514 This recorded information is shown in event displays similar to those illustrated
1515 in Figure 5.3 for simulated Monte Carlo events. To be useful for physics analyses,
1516 this series of PMT hit information needs to be reconstructed to determine the
1517 number and identity of particles and their kinematics (or track parameters): four-
1518 vertex, direction, and momentum. The reconstruction uses the fact that the charge
1519 and timing distribution of photons generated by a particular particle in an event is
1520 dependent upon its initial kinematics. Electron and muon rings are distinguished
1521 by their “fuzziness”. Muons are heavier and less affected by scattering or
1522 showering meaning they typically produce “crisp” rings. Electrons are more
1523 likely to interact via electromagnetic showering or scattering which results in
1524 larger variations of their direction from the initial direction. Consequently,
1525 electrons typically produce “fuzzier” rings compared to muons.



Figure 5.3: Event displays from Monte Carlo simulation at Super Kamiokande illustrating the “crisp” ring from a muon and the typically “fuzzy” electron ring. Each pixel represents a PMT and the color scheme denotes the accumulated charge deposited on that PMT. Figures taken from [177].

5.2 Event Reconstruction at SK

For the purposes of this analysis, the `fitQun` reconstruction algorithm [178] is utilised. Its core function is to compare a prediction of the accumulated charged and timing distribution from each PMT, generated for a particular particle identity and track parameters, to that observed in the neutrino event. It determines the preferred values by maximising a likelihood function (or minimising a log-likelihood function) which includes information from PMTs which were hit and those that were not hit. The `fitQun` algorithm is based on the key concepts of the MiniBooNE reconstruction algorithm [179].

The `fitQun` algorithm improves upon the previous `APFit` algorithm [180] which has been used for many previous SK analyses. `APFit` fits the vertex from timing information and then fits the direction of the particle from PMT hits within a 43 deg Cherenkov cone (assuming an ultra-relativistic particle) using a fitting estimator. A Hough transformation is used to find the radius of a ring (related to the momentum through Equation 3.2) as well as the number of rings contained within the event. The analysis presented here uses the `fitQun` algorithm as it improves both the accuracy of the fit parameters and the rejection of neutral

1543 current π^0 events as compared to APFit [181, 182].

1544 Any event in SK can consist of prompt (or primary) and decay (or secondary)
1545 particles. For example, a charged current muon neutrino interaction can gen-
1546 erate two particles that have the potential of generating Cherenkov photons
1547 (assuming the proton is below the Cherenkov threshold): the prompt muon,
1548 and the secondary decay-electron from the muon, approximately $2\mu\text{s}$ later. To
1549 reconstruct all particles within an event, it is divided into time clusters which are
1550 called “subevents”. Subevents after the primary subevent are considered to
1551 be decay electrons.

1552 The main steps of the `fitQun` reconstruction algorithm are:

- 1553 • **Vertex pre-fitting:** An estimate of the vertex is made using a goodness-of-fit
1554 metric based on PMT hit times
- 1555 • **Peak finding:** The initial time of each subevent is determined by clustering
1556 events by time residuals
- 1557 • **Single-ring fits:** Given the pre-fit vertex and estimated time of interaction,
1558 a maximum likelihood technique searches for a single particle generating
1559 light. Electron, muon, charged pion, and proton hypotheses are considered
- 1560 • **Multi-ring fits:** Seeded from the single-ring fits, hypotheses with multiple
1561 light-producing particles are considered using the same maximum likeli-
1562 hood technique. Electron-like or charged pion-like rings are added until
1563 the likelihood stops improving

1564 To find all the subevents in an event, a vertex goodness metric is calculated
1565 for some vertex position \vec{x} and time t ,

$$G(\vec{x}, t) = \sum_i^{\text{hit PMTs}} \exp \left(-\frac{1}{2} \left(\frac{T_{Res}^i(\vec{x}, t)}{\sigma} \right)^2 \right), \quad (5.1)$$

1566 where

$$T_{Res}^i(\vec{x}, t) = t^i - t - |R_{PMT}^i - \vec{x}| / c_n, \quad (5.2)$$

1567 is the residual hit time. It is the difference in time between the PMT hit time
 1568 t^i , of the i^{th} PMT, and the expected time of the PMT hit if the photon was at
 1569 the vertex. R_{PMT}^i is the position of the i^{th} PMT, c_n is the speed of light in water
 1570 and $\sigma = 4\text{ns}$ which is comparable to the time resolution of the PMT. When the
 1571 proposed fit values of time and vertex are close to the true values, $T_{Res}^i(\vec{x}, t)$ tends
 1572 to zero resulting in subevents appearing as spikes in the goodness metric. The
 1573 proposed fit vertex and time are grid-scanned, and the values which maximise
 1574 the goodness metric are selected as the “pre-fit vertex”. Whilst this predicts a
 1575 vertex for use in the clustering algorithm, the final vertex is fit using the higher-
 1576 precision maximum likelihood method described below.

1577 Once the pre-fit vertex has been determined, the goodness metric is scanned as
 1578 a function of t to determine the number of subevents. A peak-finding algorithm
 1579 is then used on the goodness metric, requiring the goodness metric to exceed
 1580 some threshold and drop below a reduced threshold before any subsequent
 1581 additional peaks are considered. The thresholds are set such that the rate of
 1582 false peak finding is minimised while still attaining good data to Monte Carlo
 1583 agreement. To improve performance, the pre-fit vertex for each delayed subevent
 1584 is re-calculated after PMT hits from the previous subevent are masked. This
 1585 improves the decay-electron tagging performance. Once all subevents have
 1586 been determined, the time window around each subevent is then defined by the
 1587 earliest and latest time which satisfies $-180 < T_{Res}^i < 800\text{ns}$. The subevents and
 1588 associated time windows are then used as seeds for further reconstruction.

1589 For a given subevent, the `fitQun` algorithm constructs a likelihood based on
 1590 the accumulated charge q_i and time information t_i from the i^{th} PMT,

$$L(\Gamma, \vec{\theta}) = \prod_j^{\text{unhit}} P_j(\text{unhit}|\Gamma, \vec{\theta}) \prod_i^{\text{hit}} \{1 - P_i(\text{unhit}|\Gamma, \vec{\theta})\} f_q(q_i|\Gamma, \vec{\theta}) f_t(t_i|\Gamma, \vec{\theta}). \quad (5.3)$$

1591 Where $\vec{\theta}$ defines the track parameters; vertex position, direction vector and
 1592 momenta, and Γ represents the particle hypothesis. $P_i(\text{unhit}|\Gamma, \vec{\theta})$ is the proba-
 1593 bility of the i^{th} tube to not register a hit given the track parameters and particle
 1594 hypothesis. The charge likelihood, $f_q(q_i|\Gamma, \vec{\theta})$, and time likelihood, $f_t(t_i|\Gamma, \vec{\theta})$,
 1595 represents the probability density function of observing charge q_i and time t_i on
 1596 the i^{th} PMT given the specified track parameters and particle hypothesis.

1597 The predicted charge is calculated based on contributions from both the
 1598 direct light and the scattered light. The direct light contribution is determined
 1599 based on the integration of the Cherenkov photon profile along the track. PMT
 1600 angular acceptance, water quality, and calibration measurements discussed in
 1601 subsection 3.1.2 are included to accurately predict the charge probability density
 1602 at each PMT. The scattered and reflected light is calculated in a similar way,
 1603 although it includes a scattering function that depends on the vertex of the
 1604 particle and the position of the PMT. The charge likelihood is calculated by
 1605 comparing the prediction to the observed charge in the PMT which is tuned
 1606 to the PMT simulation.

1607 The time likelihood is approximated to depend on the vertex \vec{x} , direction \vec{d} ,
 1608 and time t of the track as well as the particle hypothesis. The expected time
 1609 for PMT hits is calculated by assuming unscattered photons being emitted from
 1610 the midpoint of the track, S_{mid} ,

$$t_{\text{exp}}^i = t + S_{\text{mid}}/c + |R_{\text{PMT}}^i - \vec{x} - S_{\text{mid}}\vec{d}|/c_n, \quad (5.4)$$

1611 where c is the speed of light in a vacuum. The time likelihood is then expressed
 1612 in terms of the residual difference between the PMT hit time and the expected
 1613 hit time, $t_{\text{Res}}^i = t^i - t_{\text{exp}}^i$. The particle hypothesis and momentum also affect the
 1614 Cherenkov photon distribution. These parameters modify the shape of the time
 1615 likelihood density since in reality not all photons are emitted at the midpoint of
 1616 the track. As with the charge likelihood, the contributions from both the direct
 1617 and scattered light to the time likelihood density are calculated separately, which
 1618 are both calculated from particle gun Monte Carlo studies.

The track parameters and particle identity which maximise $L(\Gamma, \vec{\theta})$ are defined as the best-fit parameters. In practice MINUIT [183] is used to minimise the value of $-\ln L(\Gamma, \vec{\theta})$. The `fiTQun` algorithm considers an electron-like, muon-like, and charged pion-like hypothesis for events with a single final state particle, denoted “single-ring events”. The particle’s identity is determined by taking the ratio of the likelihood of each of the hypotheses. For instance, electrons and muons are distinguished by considering the value of $\ln(L(e, \vec{\theta}_e)/L(\mu, \vec{\theta}_\mu))$ in comparison to the reconstructed momentum of the electron hypothesis, as illustrated by Figure 5.4. The coefficients of the discriminator between electron-like and muon-like events are determined from Monte Carlo studies [178]. Similar distributions exist for distinguishing electron-like events from π^0 -like events, and muon-like events from pion-like events. The cuts are defined as,

$$\begin{aligned} \text{Electron/Muon} : & \ln(L_e/L_\mu) > 0.2 \times p_e^{rec} [\text{MeV}], \\ \text{Electron}/\pi^0 : & \ln(L_e/L_{\pi^0}) < 175 - 0.875 \times m_{\gamma\gamma} [\text{MeV}], \\ \text{Muon/Pion} : & \ln(L_\mu/L_{\pi^\pm}) < 0.15 \times p_\mu^{rec} [\text{MeV}], \end{aligned} \quad (5.5)$$

as taken from [184], where p_e^{rec} and p_μ^{rec} are the reconstructed momentum of the single-ring electron and muon fits, respectively. $m_{\gamma\gamma}$ represents the reconstructed invariant mass of the two photons emitted from π^0 decay. Typically, the distance between a particular entry in these two-dimensional distributions and the cut-line is termed the PID parameter and is illustrated in Figure 5.5.

The `fiTQun` algorithm also considers a π^0 hypothesis. To do this, it performs a fit looking for two standard electron-hypothesis tracks which point to the same four-vertex. This assumes the electron tracks are generated from photon-conversion so the electron tracks actually appear offset from the proposed π^0 vertex. For these fits, the conversion length, direction, and momentum of each photon are also considered as track parameters which are then fit in the same methodology as the standard single-ring hypotheses.

Whilst lower energy events are predominantly single-ring events, higher energy neutrino events can generate final states with multiple particles which

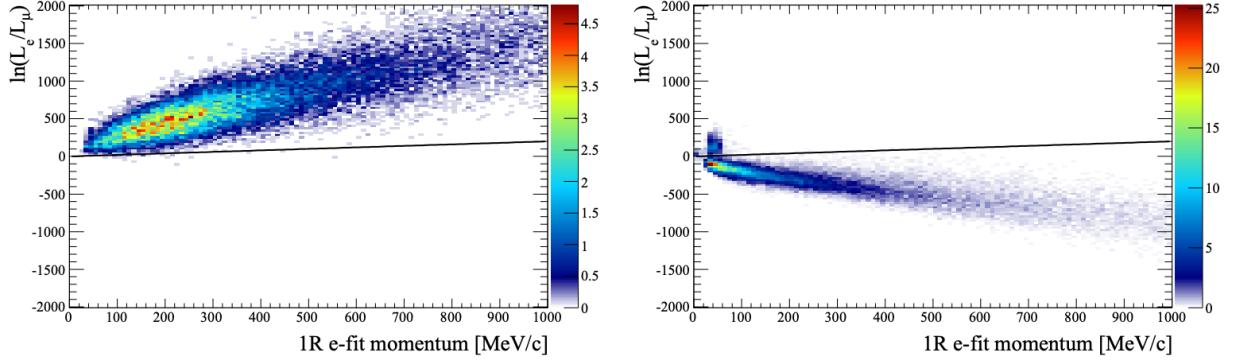


Figure 5.4: The difference of the electron-like and muon-like log-likelihood compared to the reconstructed single-ring fit momentum for atmospheric ν_e (left) and ν_μ (right) samples. The black line represents the cut used to discriminate electron-like and muon-like events, with coefficients obtained from Monte Carlo studies. Figures from [178].

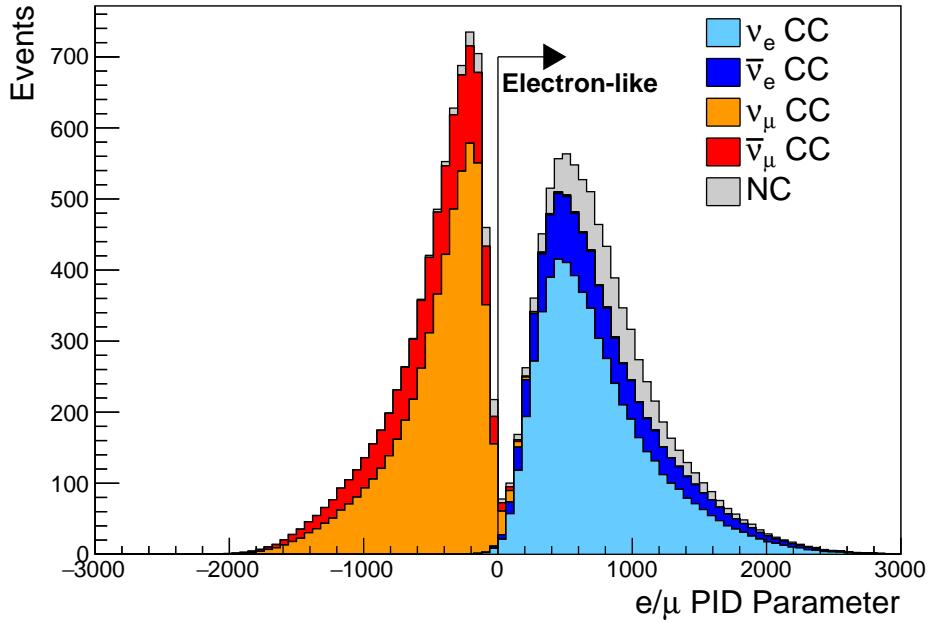


Figure 5.5: The electron/muon PID separation parameter for all sub-GeV single-ring events in SK-IV. The charged current (CC) component is broken down in four flavours of neutrino (ν_μ , $\bar{\nu}_\mu$, ν_e and $\bar{\nu}_e$). Events with positive values of the parameter are determined to be electron-like.

1645 generate Cherenkov photons. These “multi-ring” hypotheses are also considered
 1646 in the `fitQun` algorithm. When calculating the charge likelihood density, the
 1647 predicted charge associated with each ring is calculated separately and then
 1648 summed to calculate the total accumulated charge on each PMT. Similarly, the
 1649 time likelihood for the multi-ring hypothesis is calculated assuming each ring

is independent. Each track is time-ordered based on the time of flight from the center of the track to the PMT and the direct light from any ring incident on the PMT is assumed to arrive before any scattered light. To reduce computational resource usage, the multi-ring fits only consider electron-like and charged pion-like rings as the pion fit can be used as a proxy for a muon fit due to their similar mass. Due to the pions ability to interact through the strong force, they are more likely to hard-scatter. That means a single charged pion can produce multiple rings in different directions. There is an additional freedom, the fraction of kinetic energy lost in a single ring segment, which is added into the `fiTQun` pion fit to cover this difference. Pion and muon rings are indistinguishable when this fraction tends to unity.

Multi-ring fits proceed by proposing another ring to the previous fit and then fitting the parameters in the method described above. Typically, multi-ring fits have the largest likelihood because of the additional degrees of freedom introduced. A likelihood value is calculated for the n -ring and $(n + 1)$ -ring hypotheses, where the additional ring is only included if the likelihood value is above 9.35, based on Monte Carlo studies in [185].

5.2.1 Validation of Reconstruction in SK-V

Understanding how the modelling of the detector conditions and stability effects the reconstruction is critical for ensuring accurate measurements. It is important to note that the detector systematics used in the 2020 T2K-only [75] oscillation analysis are determined using data-to-Monte Carlo comparisons of the SK-IV data [186]. Due to tank-open maintenance occurring between SK-IV and SK-V, the dark rate of each PMT was observed to increase in SK-V due to light exposure for a significant time during the repairs. This increase can be seen in Figure 5.6. Run-10 of the T2K experiment was conducted in the SK-V period, so the consistency of SK-IV and SK-V data needs to be studied to determine whether the SK-IV-defined systematics can be applied to the run-10 data. Consequently, the author of this thesis assessed the quality of `fiTQun` event reconstruction for SK-V data.

This comparison study was performed using the stopping muon data set for both the SK-IV and SK-V periods. This data sample is used due to the high rate of interactions ($O(200)$ events per hour) as well as having similar energies to muons from CCQE ν_μ interactions from beam interactions. The rate of cosmic muons does depend on the solar activity cycle [187] but has been neglected in this comparison study. This is because the shape of the distributions is most important for the purposes of being compared to the detector systematics. The SK-IV and SK-V data samples consist of 2398.42 and 626.719 hours of data which equates to 686k and 192k events respectively. These samples do not correspond to the full data sets of either period but do contain enough events to be systematics limited rather than statistics limited.

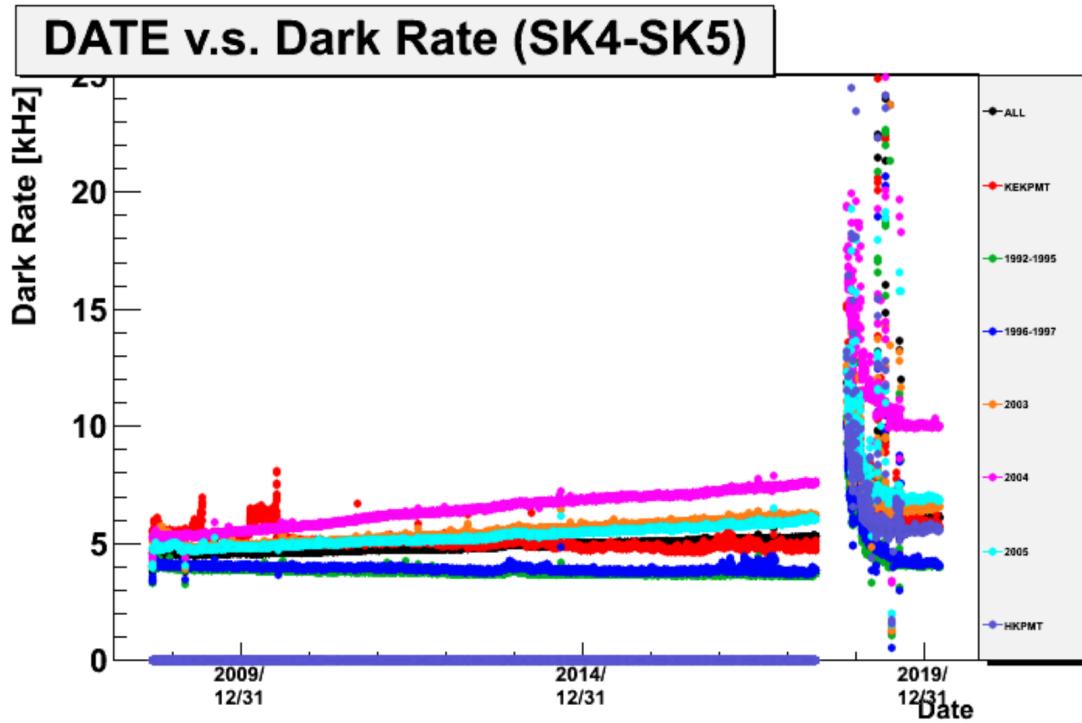


Figure 5.6: The variation of the measured dark rate as a function of date, broken down by PMT type. The SK-IV and SK-V periods span September 2008 to May 2018 and January 2019 to July 2020, respectively. The break in measurement in 2018 corresponds to the period of tank repair and refurbishment. Figure adapted from [186].

The predicted charge calculated in the `fitQun` algorithm includes a contribution from the photoelectron emission due to dark noise. Therefore, the increase

in the SK-V dark rate needs to be accounted for. In practice, the average dark rate in each SK period is calculated and used as an input in the reconstruction. This is calculated by averaging the dark rate per run for each period separately, using the calibration measurements detailed in subsection 3.1.2. The average dark rate from SK-IV and SK-V were found to be 4.57kHz and 6.30kHz, respectively. The charges associated with the muon and decay electron subevents are illustrated in Figure 5.7. The photoelectron emission from dark noise is more significant for events that have lower energy. This is because this contribution becomes more comparable to the number of photoelectrons emitted from incident photons in lower-energy events. This behaviour is observed in the data, where the charge deposited by the muon subevent is mostly unaffected by the increase in dark rate, whilst the charge associated with the decay-electron is clearly affected.

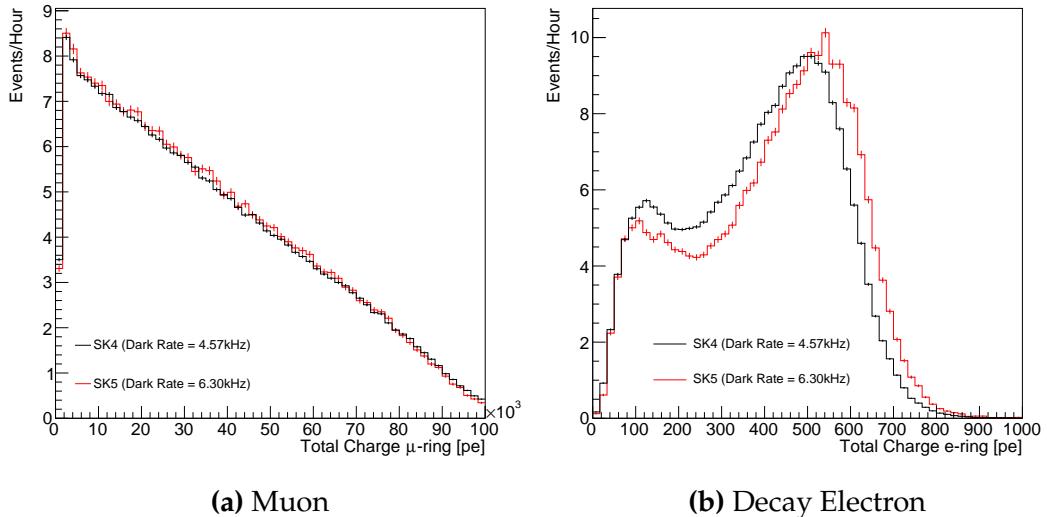


Figure 5.7: Comparison of the measured raw charge deposited per event from the stopping muon data samples between SK-IV (Blue) and SK-V (Red), split by the primary muon subevent (left) and the associated decay electron subevent (right).

The energy scale systematic is estimated from data-to-Monte Carlo differences in the stopping muon sample in [188] and found to be 2.1%. To determine the consistency of SK-IV and SK-V with respect to the energy scale systematic, the muon momentum distribution is compared between the two SK periods. As the total number of Cherenkov photons is integrated across the track length,

₁₇₀₉ the reconstructed momentum divided by track length (or range) is compared
₁₇₁₀ between SK-IV and SK-V as illustrated in Figure 5.8.

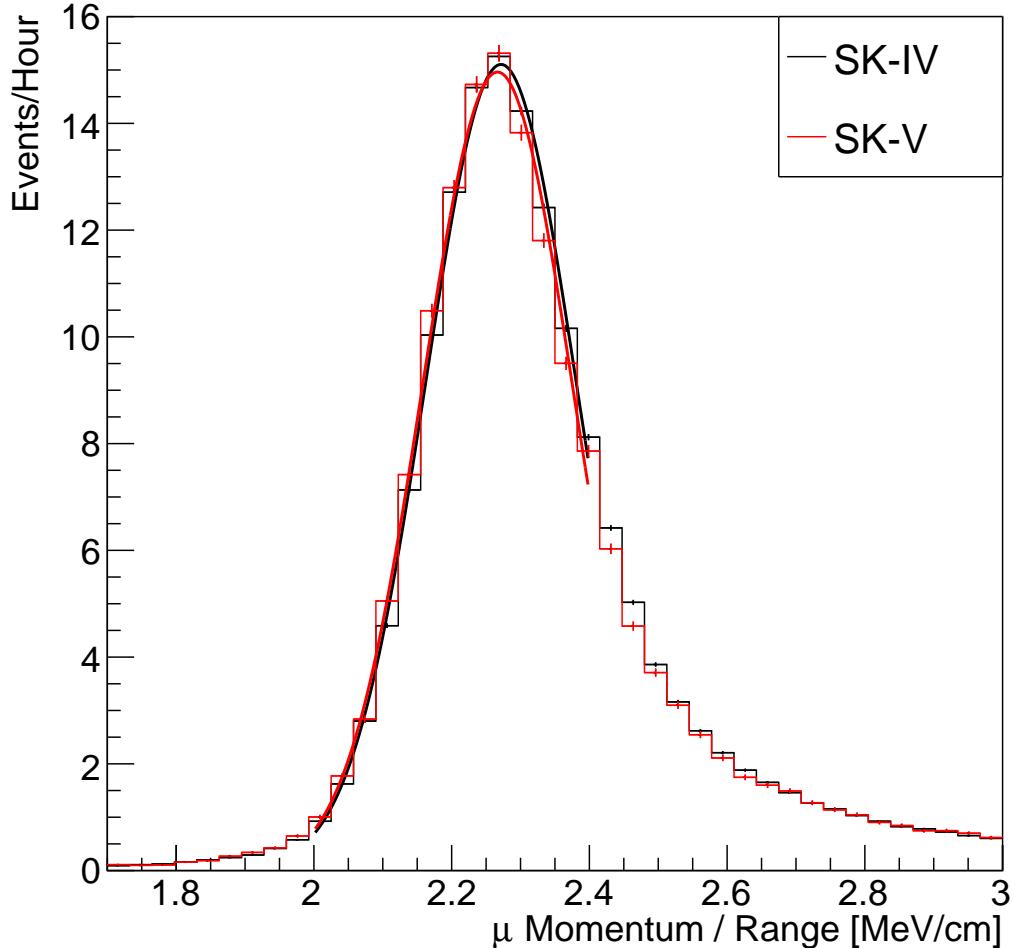


Figure 5.8: The distribution of the reconstructed momentum from the muon ring divided by the distance between the reconstructed muon and decay electron vertices as found in the stopping muon data sets of SK-IV (Black) and SK-IV (Red). Only events with one tagged decay electron are considered. A Gaussian fit is considered in the range $[2.0, 2.4]\text{MeV}/\text{cm}$ and illustrated as the solid curve.

₁₇₁₁ The consistency between these muon distributions has been computed in two
₁₇₁₂ ways. Firstly, a Gaussian is fit to the peak of each distribution separately, whose
₁₇₁₃ mean is found to be $(2.272 \pm 0.003)\text{MeV}/\text{cm}$ and $(2.267 \pm 0.006)\text{MeV}/\text{cm}$ for SK-
₁₇₁₄ IV and SK-V respectively. The ratio of these is equal to 1.002 ± 0.003 . The means of
₁₇₁₅ the Gaussian fits are consistent with the expected stopping power of a minimum

¹⁷¹⁶ ionising muon for a target material (water) with $Z/A \sim 0.5$ [189]. The second
¹⁷¹⁷ consistency check is performed by introducing a nuisance parameter, α , which
¹⁷¹⁸ modifies the SK-V distribution. The value of α which minimises the χ^2 value
¹⁷¹⁹ between the SK-IV and SK-V is determined by scanning across a range of values.
¹⁷²⁰ This is repeated by applying the nuisance parameter as both a multiplicative
¹⁷²¹ factor and an additive shift. The χ^2 distributions for different values of α is
¹⁷²² illustrated in Figure 5.9. The values which minimise the χ^2 are found to be 0.0052
¹⁷²³ and 1.0024 for the additive and multiplicative implementations, respectively. No
¹⁷²⁴ evidence of shifts larger than the 2.1% uncertainty on the energy scale systematic
¹⁷²⁵ has been found in the reconstructed momentum distribution of SK-IV and SK-V.

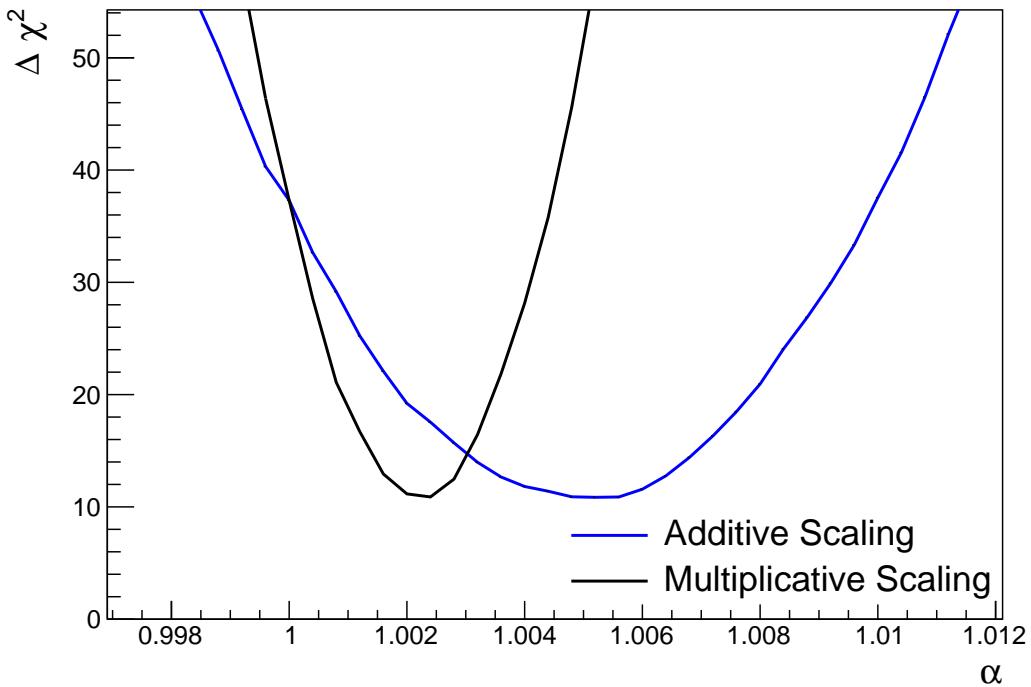


Figure 5.9: The χ^2 difference between the SK-IV and SK-V reconstructed muon momentum divided by range when the SK-V is modified by the scaling parameter α . Both additive (Blue) and multiplicative (Black) scaling factors have been considered. In practice, the additive scaling factor actually uses the value of $(\alpha - 1.0)$ but is illustrated like this so the results can be shown on the same axis range.

1726 5.3 Event Reduction at SK

1727 In normal data-taking operations, the SK detector observes many background
 1728 events alongside the beam and atmospheric neutrino signal events of physics
 1729 interest for this thesis. Cosmic ray muons and flasher events, which are the spon-
 1730 taneous discharge of a given PMT, contribute the largest amount of background
 1731 events in the energy range relevant to this thesis. Therefore the data recorded
 1732 is reduced with the aim of removing these background events. The reduction
 1733 process is detailed in [55, 88] and briefly summarised below.

1734 Atmospheric neutrino events observed in the SK detector are categorised
 1735 into three different types of samples: fully contained (FC), partially contained
 1736 (PC) and up-going muon (Up- μ), using PMT hit signatures in the inner and
 1737 outer detector (ID and OD, respectively). To identify FC neutrino events, it is
 1738 required that the neutrino interacts inside the fiducial volume of the ID and that
 1739 no significant OD activity is observed. For this analysis, an event is defined to be
 1740 in the fiducial volume provided the event vertex is at least 0.5m away from the
 1741 ID walls. PC events have the same ID requirements but can have a larger signal
 1742 present inside the OD. Typically, only high energy muons from ν_μ interactions can
 1743 penetrate the ID wall. The Up- μ sample contains events where muons are created
 1744 from neutrino interactions in the OD water or rock below the tank. They then
 1745 propagate upwards through the detector. Downward-going muons generated
 1746 from neutrino interactions above the tank are neglected because of the difficulty
 1747 in separating their signature from the cosmic muon shower background. The
 1748 sample categories are visually depicted in Figure 5.10.

1749 Based on the event characteristics, as defined by the `fitQun` event reconstruc-
 1750 tion software, the FC events are categorised by

- 1751 • **Visible Energy:** equal to the sum of the reconstructed kinetic energy of
 1752 particles above the Cerenkov threshold for all rings present in the event.
 1753 The purpose is to separate events into sub-GeV and multi-GeV categories.

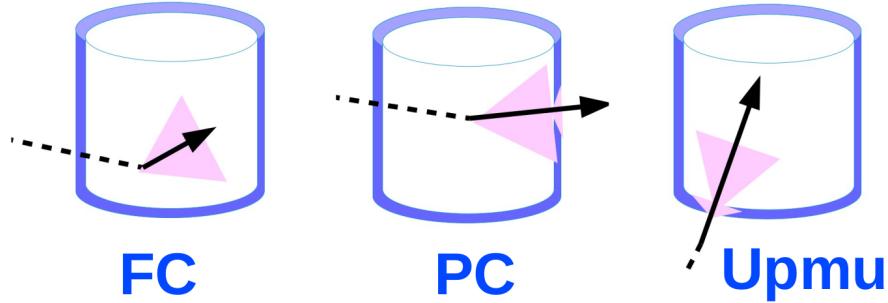


Figure 5.10: A depiction of the topology patterns for fully-contained (FC), partially-contained (PC), and up-going muon ($\text{Up-}\mu$) samples included in this analysis.

- **Number of observed Cerenkov rings.** The purpose is to separate single-ring and multi-ring events, where single-ring events predominantly consist of quasi-elastic interactions and multi-ring events are typically resonant pion production or deep inelastic scattering events.
- **Particle identification parameter of the most energetic ring:** A value determined from the maximum likelihood value based on `fitQun`'s electron, muon, or pion hypothesis. The purpose is to separate electron-like and muon-like events.
- **Number of decay electrons:** The purpose is to separate quasi-elastic events (which have one decay electron emitted from the muon decay) and resonant pion production events (which have two decay electrons emitted from the muon and pion).

The PC and Up- μ categories are broken down into “through-going” and “stopping” samples depending on whether the muon leaves the detector. This is because the PC stopping events deposit the entire energy of the interaction into the detector, resulting in better reconstruction. The energy of events that exit the detector has to be estimated, with a typically worse resolution, which introduces much larger systematic uncertainties. Through-going Up- μ samples are further broken down by whether any hadronic showering was observed in the event which typically indicates DIS interactions. The expected neutrino energy for the different categories is given in Figure 5.11. FC sub-GeV and multi-GeV events

¹⁷⁷⁵ peak around 0.7GeV and 3GeV respectively, with slightly different peak energies
¹⁷⁷⁶ for ν_e and ν_μ oscillation channels. PC and Up- μ are almost entirely comprised
¹⁷⁷⁷ of ν_μ events and peak around 7GeV and 100GeV, respectively.

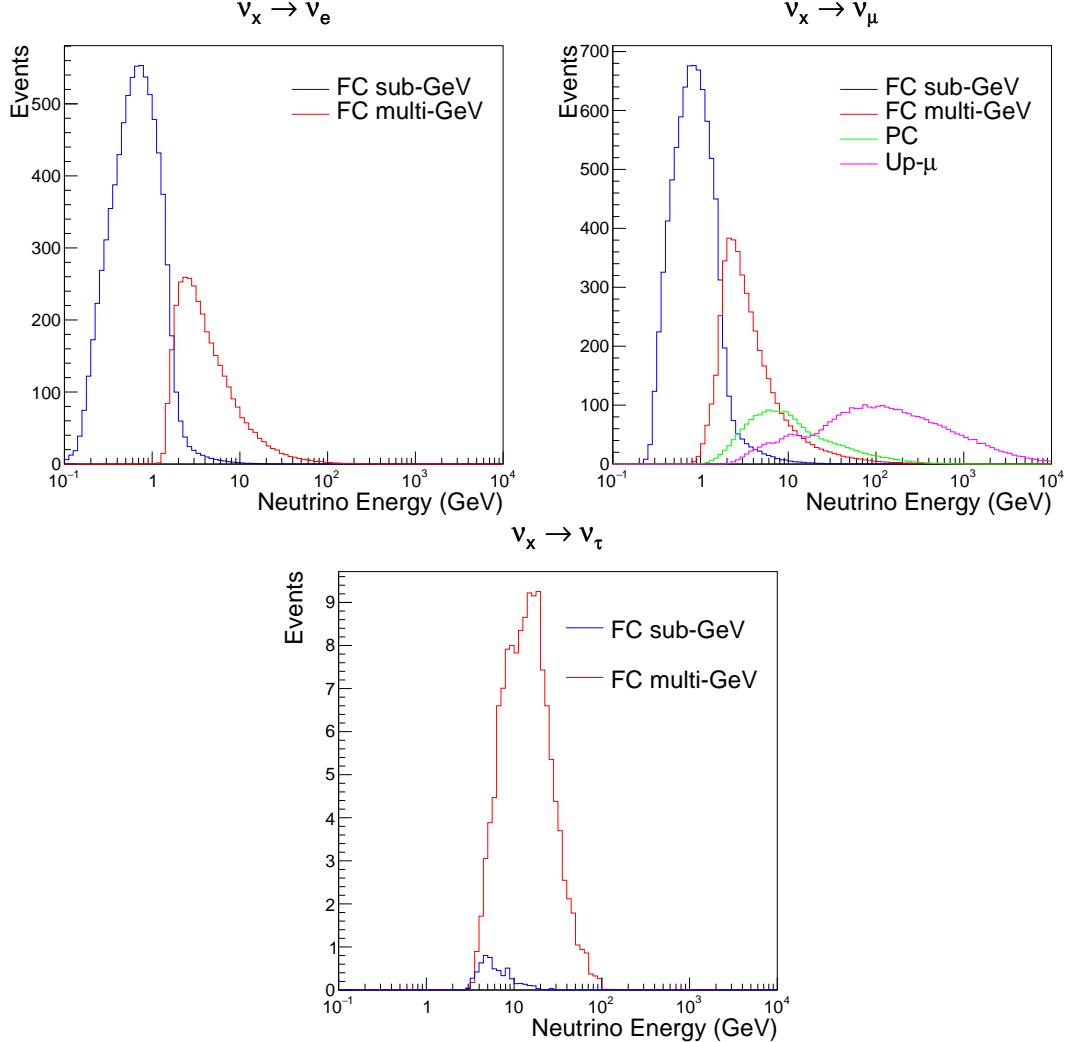


Figure 5.11: The predicted neutrino flux of the fully contained (FC) sub-GeV and multi-GeV, partially contained (PC), and upward-going muon (Up- μ) events. The prediction is broken down by the $\nu_x \rightarrow \nu_e$ prediction (top left), $\nu_x \rightarrow \nu_\mu$ prediction (top right) and $\nu_x \rightarrow \nu_\tau$ prediction (bottom). ν_x represents the flavours of neutrinos produced in the cosmic ray showers (electron and muon). Asimov A oscillation parameters are assumed (given in Table 2.2).

¹⁷⁷⁸ The first two steps in the FC reconstruction remove the majority of cosmic
¹⁷⁷⁹ ray muons by requiring a significant amount of ID activity compared to that
¹⁷⁸⁰ measured in the OD. Events that pass this cut are typically very high momentum
¹⁷⁸¹ muons or events that leave very little activity in the OD. Consequently, a third

reduction step is then applied to select cosmic-ray muons that pass the initial reduction step. A purpose-built cosmic muon fitter is used to determine the entrance (or exit) position of the muon and a cut is applied to OD activity contained within 8m of this position. Flasher events are removed in the fourth reduction step which is based on the close proximity of PMT hits surrounding the PMT producing the flash. Events that pass all these reduction steps are reconstructed with the APFit algorithm. The fifth step of the reduction uses information from the more precise fitter to repeat the previous two steps with tighter cuts. Muons below the Cherenkov threshold can not generate optical photons in the ID but the associated decay electron can due to its lower mass. These are the types of events targeted in the fifth reduction step. The final cuts require the event vertex to be within the fiducial volume (0.5m from the wall although the nominal distance is 2.0m), visible energy $E_{vis} > 30\text{MeV}$ and fewer than 16 hits within the higher energy OD cluster. The culmination of the fully contained reduction results in 8.09 events/day in the nominal fiducial volume [84]. The uncertainty in the reconstruction is calculated by comparing Monte Carlo prediction to data. The largest discrepancy is found to be 1.3% in the fourth reduction step.

The PC and Up- μ events are processed through their own reduction processes detailed in [55]. Both of these samples are reconstructed with the APFit algorithm rather than `f1TQun`. This is because the efficiency of reconstructing events that leave the detector has not been sufficiently studied for reliable systematic uncertainties with `f1TQun`. The PC and Up- μ samples acquire events at approximately 0.66 and 1.44 events/day.

Beam neutrinos events undergo the same reduction steps as FC events and are then subject to further cuts [190]. The GPS system that links the timing between the beam facility and SK needs to be operating correctly and there should be no activity within the detector in the previous $100\mu\text{s}$ before the trigger. The events then need to triggered between $-2\mu\text{s}$ and $10\mu\text{s}$ of the expected spill timing.

1811 The beam neutrino samples are not split by visible energy since their energy
 1812 range is smaller than the atmospheric neutrino events. Following the T2K
 1813 analysis in [75], only single-ring beam neutrino events are considered. Similar to
 1814 atmospheric event selection, the number of decay electrons is used as a proxy for
 1815 distinguishing CCQE and CCRES events. The expected neutrino energy, broken
 1816 down by the number of decay electrons, is given in Figure 5.12.

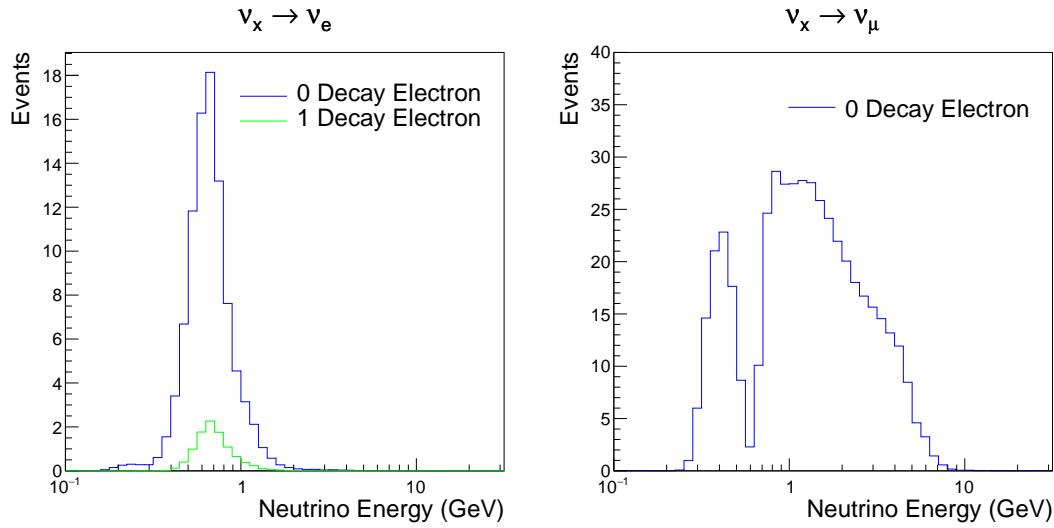


Figure 5.12: The predicted flux of beam neutrinos, as a function of neutrino energy. The predictions are broken down by the number of decay electrons associated with the particular events. Asimov A oscillation parameters are assumed (given in Table 2.2).

6

1817

1818

Sample Selections and Systematics

1819 The oscillation analysis presented within this thesis is built upon a simultaneous
1820 fit to atmospheric samples at SK, neutrino beam samples in the near detector,
1821 and beam samples at SK. This is the first simultaneous oscillation analysis of
1822 beam and atmospheric samples supported by the T2K and SK collaborations.
1823 Notably, the author of this thesis has been responsible for the building and
1824 developing the MaCh3 framework to support all sets of samples simultaneously.
1825 The definitions of the samples are documented in section 6.1, section 6.2, and
1826 section 6.3, respectively. The data collected and used within this analysis is
1827 detailed in Table 6.1. The near and far detector data corresponds to T2K runs
1828 2-9 and runs 1-10, respectively. The accumulated POT and beam power for runs
1829 1 – 10 are illustrated in Figure 6.1.

Data Type	Total
Near Detector FHC	1.15×10^{21} POT
Near Detector RHC	8.34×10^{20} POT
Far Detector FHC	1.97×10^{21} POT
Far Detector RHC	1.63×10^{21} POT
Atmospheric SK-IV	3244.4 days

Table 6.1: The amount of data collected in each detector used within this analysis. The data collected at the near and far detector, for both neutrino beam (FHC) and antineutrino beam (RHC), is measured as the number of protons on target (POT).

1830 The difference in POT recorded at the near and far detector is due to the
 1831 difference in downtime. The SK detector is very stable with almost 100% of
 1832 data recorded during beam operation. Due to various technical and operational
 1833 issues, the downtime of the near detector is significantly higher due to its more
 1834 complex design and operating requirements.

1835 The systematic parameters invoked within the flux, detector, and interaction
 1836 models used within this analysis are documented in section 6.4. The standard
 1837 configuration of the joint beam and atmospheric data fit utilises far detector sys-
 1838 tematics provided in the official inputs from the two experiments. Additionally,
 1839 a correlated detector model which fits the parameters used in sample selections
 1840 to data has been developed and documented in subsection 6.4.5.

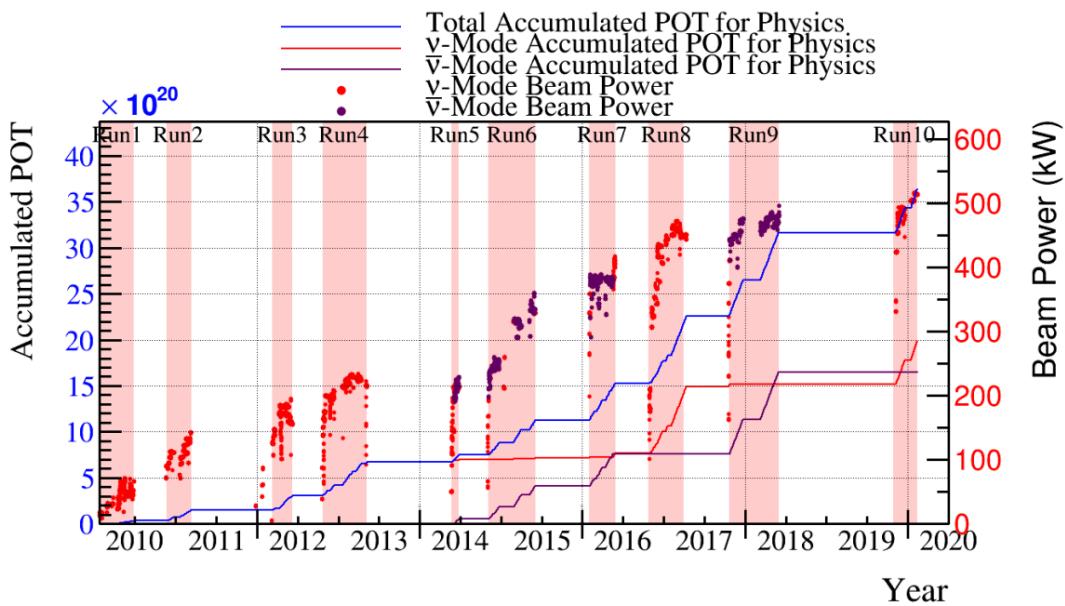


Figure 6.1: The accumulated beam data, measured as the number of protons on target (POT). The total data (blue) is given which comprises of the neutrino beam (red) and antineutrino (purple) components. The beam power for neutrino and antineutrino beams is given as the markers using the same colour scheme. The timescale runs from Run 1 which started in January 2010 until Run 10 which ended in February 2020. The ratio of accumulated data in neutrino and antineutrino beam is 54.7% : 45.3%.

1841 6.1 Atmospheric Samples

1842 The atmospheric event selection follows the official SK-IV analysis presented
1843 in [88] and is documented below. The Monte Carlo prediction used within this
1844 analysis corresponds to 500 years worth of neutrino events, which is scaled down
1845 to match the SK-IV livetime of 3244.4 days.

1846 The fully contained (FC), partially contained (PC), and upward going muon
1847 events ($\text{up-}\mu$) which pass the reduction cuts discussed in section 5.3 are further
1848 broken down into different samples based on reconstruction information. This
1849 section details the samples used within this oscillation analysis, alongside the
1850 chosen binning.

1851 FC events are first separated by the visible energy deposited within the
1852 detector. This is calculated as the sum of the reconstructed kinetic energy
1853 above the Cherenkov threshold for all rings present in the event. Events are
1854 separated by whether they were above or below $E_{\text{vis}} = 1.33\text{GeV}$. This separates
1855 “subGeV” and “multiGeV” events. Typically, lower energy events consist of
1856 charged current quasi-elastic (CCQE) interactions which are better understood
1857 and simpler to reconstruct resulting in smaller systematic uncertainties. Events
1858 are further separated by the number of rings associated with the event due to
1859 similar reasoning. As the oscillation probability is dependant upon the flavour
1860 of neutrino, electron and muon events are separated using a similar likelihood
1861 method to that discussed in section 5.2. To reduce computational resources
1862 required for the reconstruction, only electron and pion hypotheses are considered
1863 so this separation cut depends on the ratio of the electron to pion likelihoods,
1864 $\log(L_e/L_\pi)$. Finally, the number of decay electrons is used to classify events.
1865 Charged current resonant pion production (CCRES) interactions generate a final-
1866 state pion. This can decay, mostly likely through a muon, into a decay electron.
1867 Therefore any electron-like event with one decay electron or muon-like event
1868 with two decay electrons was most likely produced by a CCRES interaction.
1869 Consequently, the number of decay electrons can be used to distinguish CCQE

1870 and CCRES interaction modes. Ultimately, FC subGeV events are separated
1871 into the samples listed in Table 6.2.

Sample Name	Description
SubGeV- <i>e</i> like-0dcy	Single ring <i>e</i> -like events with zero decay electrons
SubGeV- <i>e</i> like-1dcy	Single ring <i>e</i> -like events with one or more decay electrons
SubGeV- <i>μ</i> like-0dcy	Single ring <i>μ</i> -like events with zero decay electrons
SubGeV- <i>μ</i> like-1dcy	Single ring <i>μ</i> -like events with one decay electrons
SubGeV- <i>μ</i> like-2dcy	Single ring <i>μ</i> -like events with two or more decay electrons
SubGeV- <i>π</i> 0like	Two <i>e</i> -like ring events with zero decay electrons and reconstructed π^0 mass $85 \leq m_{\pi^0} < 215$ MeV

Table 6.2: The fully contained subGeV samples, defined as events with visible energy $E_{vis} < 1.33$ GeV, used within this oscillation analysis.

1872 In addition to the cuts discussed above, multiGeV samples also have addi-
1873 tional cuts to separate samples which target neutrino and antineutrino events.
1874 As discussed in section 2.5, the matter resonance only occurs for neutrinos in the
1875 normal hierarchy and antineutrinos in the inverted mass hierarchy. Therefore,
1876 having flavour-enriched samples aids in the determination of the mass hierarchy.
1877 For a CCRES interaction,

$$\begin{aligned}
 \bar{\nu}_e + N &\rightarrow e^+ + N' + \pi^-, \\
 \nu_e + N &\rightarrow e^- + N' + \pi^+ \\
 &\quad \downarrow \mu^+ + \nu_\mu \\
 &\quad \downarrow e^+ + \nu_e + \bar{\nu}_\mu.
 \end{aligned} \tag{6.1}$$

1878 The π^- emitted from a $\bar{\nu}_e$ interaction is more likely to be captured by an
1879 oxygen nucleus than the π^+ from ν_e interactions [191]. These pions then decay,
1880 mostly through muons, to electrons. Therefore the number of tagged decay
1881 electrons associated with an event gives an indication of whether the interaction
1882 was due to a neutrino or antineutrino: zero for $\bar{\nu}_e$ events, and one for ν_e events.
1883 The ability to separate neutrino from antineutrino events is illustrated in Table 6.4,
1884 where the MultiGeV-*e*like-nue has 78% purity of CC neutrino interactions with
1885 only 7% antineutrino background, the rest consisting of NC backgrounds.

1886 The number of decay electrons discriminator works reasonably well for single-
 1887 ring events. However, this is not the case for multi-ring events. A multiGeV
 1888 multiring electron-like (MME) likelihood cut was introduced in [192, 193]. This
 1889 is a two-stage likelihood selection cut. Four observables are used in the first
 1890 likelihood cut to distinguish $CC\nu_e$ and $CC\bar{\nu}_e$ events from background:

- 1891 • The number of decay electrons
 1892 • The maximum distance between the vertex of the neutrino and the decay
 1893 electrons
 1894 • The energy deposited by the highest energy ring
 1895 • The particle identification of that highest energy ring

1896 Background events consist of $CC\nu_\mu$ and NC interactions. Typically, the
 1897 majority of the energy in these background events is carried by the hadronic
 1898 system. Additionally, muons tend to travel further than the pions from $CC\nu_e$
 1899 before decaying. Thus, the parameters used within the likelihood cut target these
 1900 typical background interaction kinematics.

Sample Name	Description
MultiGeV-elike-nue	Single ring e -like events with zero decay electrons
MultiGeV-elike-nuebar	Single ring e -like events with one or more decay electrons
MultiGeV-mulike	Single ring μ -like events
MultiRing-elike-nue	Two or more ring events with leading energy e -like ring and passed both MME and $\nu/\bar{\nu}$ separation cuts
MultiRing-elike-nuebar	Two or more ring events with leading energy e -like ring and passed MME and failed $\nu/\bar{\nu}$ separation cuts
MultiRing-mulike	Two or more ring events with leading energy μ -like ring and only requires $E_{vis} > 0.6\text{GeV}$
MultiRing-Other1	Two or more ring events with leading energy e -like ring and failed the MME likelihood cut

Table 6.3: The fully contained multiGeV samples used within this oscillation analysis. Both the sample name and description are given.

1901 Neutrino and antineutrino events are then separated by a second likelihood
 1902 method ($\nu/\bar{\nu}$ separation) detailed in [60]. This uses the number of decay electrons,

1903 the number of reconstructed rings, and the event’s transverse momentum. The
1904 last two parameters are used because higher-energy samples tend to have more
1905 pions produced above the Cherenkov threshold which results in more rings
1906 compared to an antineutrino interaction. Furthermore, the angular distribution
1907 also tends to be more forward peaked in antineutrino interactions as compared
1908 to neutrino interactions [88]. These FC multiGeV sample definitions are de-
1909 tailed in Table 6.3.

1910 The PC and up- μ samples are split by the amount of energy deposited within
1911 the outer detector, into “stopping” and “through-going” samples. If an event
1912 leaves the detector, the energy it takes with it has to be estimated which increases
1913 the systematic uncertainty compared to events entirely contained within the
1914 inner detector. This estimation is particularly poor at high energies, thus the
1915 up- μ through-going events are not binned in reconstructed momentum. The
1916 through-going up- μ are further separated by the presence of any electromagnetic
1917 showering in the event, as the assumption of non-showering muon does not give
1918 reliable reconstruction for these types of events [55]. In total, 13 FC, 2 PC, and
1919 3 up- μ atmospheric samples are included within this analysis.

1920 The atmospheric samples are binned in direct observables: reconstructed
1921 lepton momentum and direction, as given by Table A.1. The distribution of
1922 the reconstructed lepton momentum (for samples that only have one bin in
1923 reconstructed zenith angle) and reconstructed direction for each atmospheric
1924 sample used within this analysis is illustrated in Figure 6.2. The by-mode
1925 breakdown of each of the atmospheric samples is given in Appendix A.

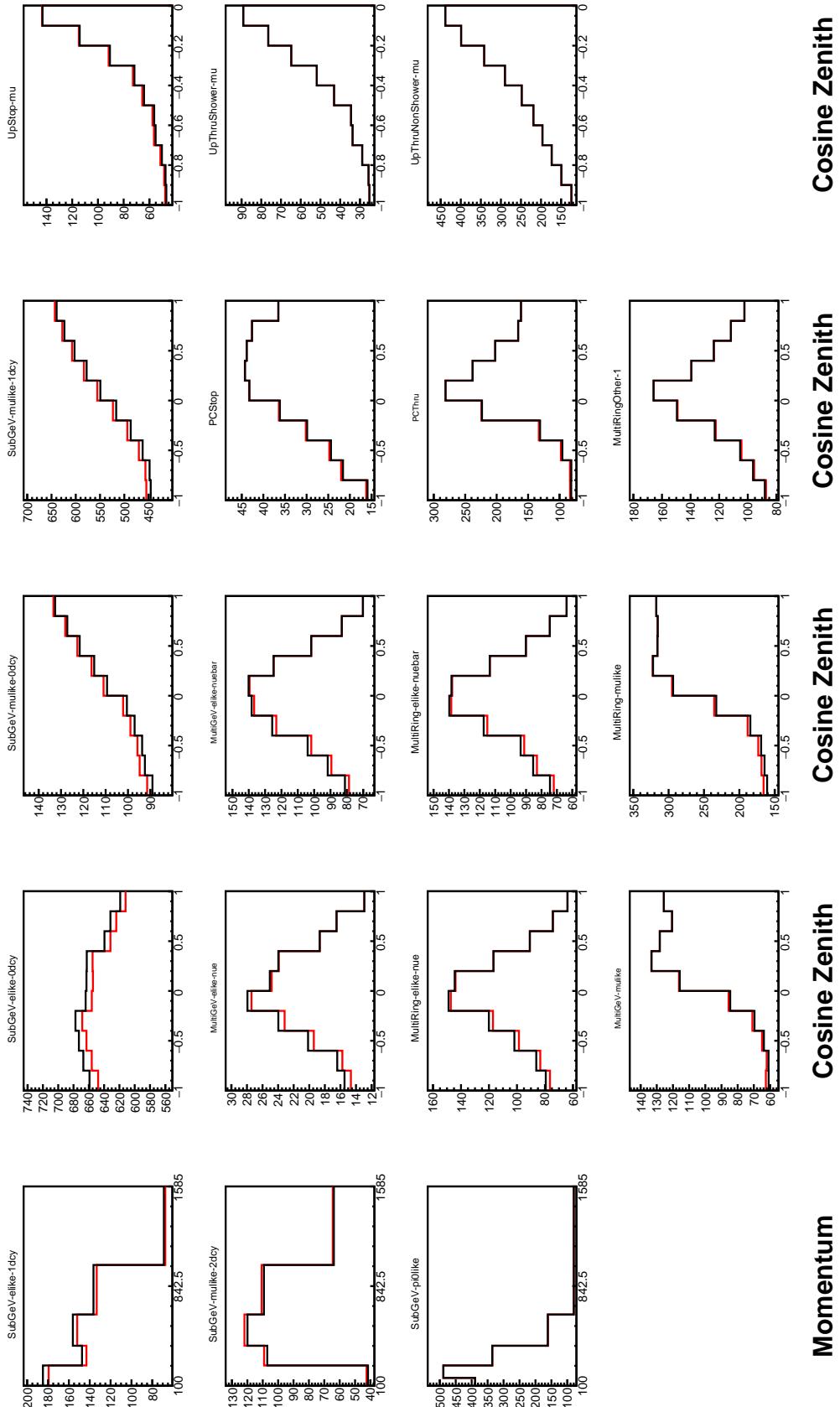


Figure 6.2: Comparison of the SK-IV atmospheric samples between predictions made with the CP-violating Asimov A (Black) and CP-conserving Asimov B (Red) oscillation parameter sets (given in Table 2.2). The subGeV samples CCRES and π^0 -like samples are given in their reconstructed lepton momentum. All other samples are presented in their reconstructed zenith angle projection.

Sample	$CC\nu_e$	$CC\bar{\nu}_e$	$CC(\nu_\mu + \bar{\nu}_\mu)$	$CC(\nu_\tau + \bar{\nu}_\tau)$	NC
SubGeV- <i>elike</i> -0dcy	72.17	23.3	0.724	0.033	3.77
SubGeV- <i>elike</i> -1dcy	86.81	1.773	7.002	0.062	4.351
SubGeV- <i>mulike</i> -0dcy	1.003	0.380	90.07	0.036	8.511
SubGeV- <i>mulike</i> -1dcy	0.023	0.	98.46	0.029	1.484
SubGeV- <i>mulike</i> -2dcy	0.012	0.	99.25	0.030	0.711
SubGeV- <i>pi0like</i>	6.923	2.368	0.928	0.011	89.77
MultiGeV- <i>elike</i> -nue	78.18	7.041	3.439	1.886	9.451
MultiGeV- <i>elike</i> -nuebar	56.68	37.81	0.174	0.614	4.718
MultiGeV- <i>mulike</i>	0.024	0.005	99.67	0.245	0.058
MultiRing- <i>elike</i> -nue	59.32	12.39	4.906	3.385	20
MultiRing- <i>elike</i> -nuebar	52.39	31.03	1.854	1.585	13.14
MultiRing- <i>mulike</i>	0.673	0.080	97.33	0.342	1.578
MultiRingOther-1	27.98	2.366	34.93	4.946	29.78
PCStop	8.216	3.118	84.45	0.	4.214
PCThru	0.564	0.207	98.65	0.	0.576
UpStop-mu	0.829	0.370	98.51	0.	0.289
UpThruNonShower-mu	0.206	0.073	99.62	0.	0.103
UpThruShower-mu	0.128	0.054	99.69	0.	0.132

Table 6.4: The purity of each atmospheric sample used within this analysis, broken down by charged current (CC) and neutral current (NC) interactions and which neutrino flavour interacted within the detector. Each row sums to 100% by definition. Asimov A oscillation parameter sets are assumed (given in Table 2.2). Electron neutrino and antineutrino events are separated to illustrate the ability of the separation likelihood cuts used within the multiGeV and multiring sample selections.

1926 6.2 Near Detector Beam Samples

1927 The near detector sample selections are documented in detail within [194] and
 1928 summarised below. Samples are selected based upon which of the two Fine
 1929 Grained Detector (FGD) the vertex is reconstructed in as well as the operating
 1930 mode of the beam: FHC or RHC. Wrong-sign neutrino background samples are
 1931 considered in the RHC mode in order to add additional constraints on model
 1932 parameters. Samples from the wrong-sign component of the FHC beam mode
 1933 are not included as they are statistically insignificant compared to those samples
 1934 already listed.

1935 The reconstruction algorithm uses a clustering algorithm to group hits within
 1936 the TPC. It then adds information from the upstream FGD to form a track
 1937 that passes through both sub-detectors. In FHC(RHC), the highest momentum
 1938 negative(positive) curvature track is defined as the muon candidate. Before
 1939 being assigned a sample, these candidate muon events must pass CC-inclusive
 1940 cuts, as defined in [195]:

- 1941 • Event Timing: The DAQ must be operational and the event must occur
 1942 within the expected beam time window consistent with the beam spill
- 1943 • TPC Requirement: The muon-candidate track path must intercept one or
 1944 more TPCs
- 1945 • Fiducial volume: The event must originate from within the fiducial volume
 1946 defined in [196]
- 1947 • Upstream Background: Remove events that have muon tracks that originate
 1948 upstream of the FGDs by requiring no high-momentum tracks within
 1949 150mm upstream of the candidate vertex. Additionally, events that occur
 1950 within the downstream FGD are vetoed if a secondary track starts within
 1951 the upstream FGD

- 1952 • Broken track removal: All candidates where the muon candidate is broken
1953 in two are removed
- 1954 • Muon PID: Measurements of dE/dx in a TPC are used to distinguish muon-
1955 like events, from electron-like or proton-like, using a likelihood cut
- 1956 In addition to these cuts, RHC neutrino events also have to undergo the
1957 following cuts to aid in the separation of neutrino and antineutrino [197]:
- 1958 • TPC Requirement: The track path must intercept TPC2
- 1959 • Positive Track: The highest momentum track must have a positive recon-
1960 structed charge
- 1961 • TPC1 Veto: Remove any events originating upstream of TPC1
- 1962 Once all CC-inclusive events have been determined, they are further split by
1963 pion multiplicity: CC0 π , CC1 π , and CCOther. Pions in the TPCs are selected by
1964 requiring a second track to be observed, which is separate from the muon track
1965 and is in the same beam spill window and sub-detector. The number of FGD
1966 pions is equal to the number of Michel electrons which were tagged within the
1967 same sub-detector and spill window. If this value is equal to zero, the number
1968 of FGD pions is equivalent to the number of pion-like tracks which have dE/dx
1969 measurements consistent with the pion hypothesis. The pion tracks from both
1970 FGD and TPC events are required to have a vertex consistent with that of the
1971 muon candidate. The Michel electron tagging is preferential as a delayed Michel
1972 is almost always a pion meaning this cut has a higher purity [195, 198], whereas a
1973 track in the FGD that is consistent with a pion could be another particle resulting
1974 in a lower purity. Michel electrons are neglected in the TPC as the pions very
1975 rarely stop due to the low density.
- 1976 CC0 π , CC1 π , and CCOther samples are defined with the following cuts:

- 1977 • ν_μ CC0 π Selection: No electrons in TPC and no charged pions or decay
1978 electrons within the TPC or FGD

- $\nu_\mu \text{CC}1\pi$ **Selection:** Exactly one charged pion in either the TPC or FGD
- $\nu_\mu \text{CCOther}$ **Selection:** All events which are not classified into the above two selections

Counting the three selections for each FGD in FHC and RHC running, including the wrong-sign background in RHC, 18 near detector samples are used within this analysis. These samples are binned in reconstructed lepton momentum (illustrated in Figure 6.3) and direction with respect to the beam. The binning is chosen such that each event has at least 20 Monte Carlo events in each bin [196]. This is to ensure that the bins are coarse enough to ensure the reduction of statistical errors, whilst also being fine enough to sample the high-resolution peak regions. The exact binning is detailed in [196].

6.3 Far Detector Beam Samples

The beam neutrino events which occur at the SK detector, which pass the reduction cuts detailed in section 5.3, are separated based on whether the beam was operating in FHC or RHC mode. The events are then separated into three samples: electron-like ($1R e$), muon-like ($1R \mu$), and $\text{CC}1\pi^+$ -like ($1R e1de$) which are observed as electron-like events with an associated decay electron [186]. As discussed in section 6.1, positively charged pions emitted from neutrino interactions are more likely to produce decay electrons than negatively charged pions. Consequently, the $\text{CC}1\pi^+$ -like sample is only selected when the beam is operating in FHC mode. Therefore, five beam samples measured at SK are used in this analysis.

The fiducial volume definition for beam samples is slightly different from that used for the atmospheric samples. It uses both the distance to the closest wall (`dWall`) and the distance to the wall along the trajectory of the particle (`toWall`). This allows events that originate close to the wall but are facing into the tank to be included within the analysis, which would have otherwise been removed. These additional events are beneficial for a statistics-limited experiment. The exact

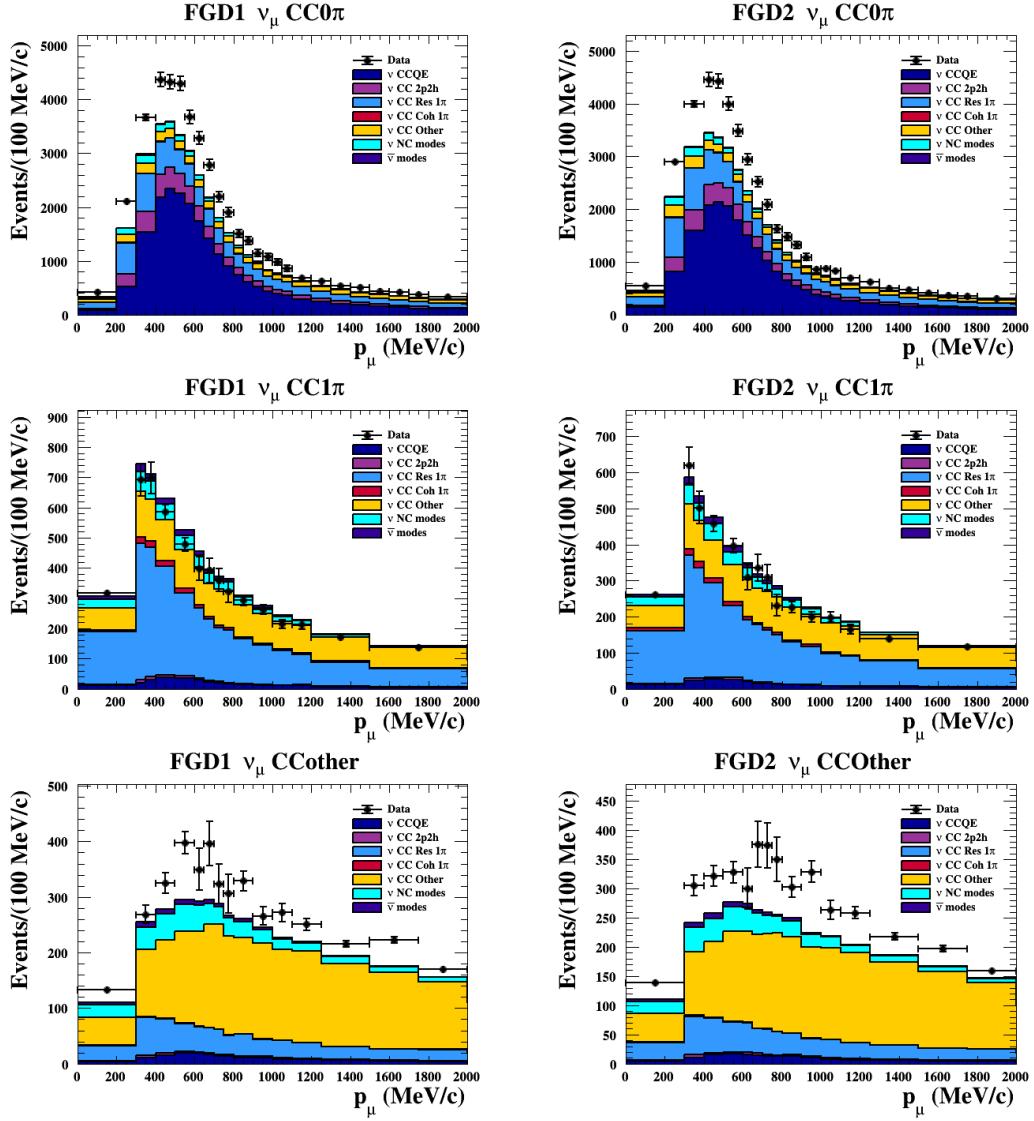


Figure 6.3: The nominal Monte Carlo predictions compared to data for the FGD1 and FGD2 samples in neutrino beam mode, broken down into the $CC\nu_\mu 0\pi$, $CC\nu_\mu 1\pi$ and $CC\nu_\mu$ Other categories. Figures taken from [194].

2007 cut values for both `dWall` and `tWall` are different for each of the three types of
2008 sample and are optimised based on T2K sensitivity to δ_{CP} [184, 199]. They are:

2009 **1Re event selection** For an event to be classified as a 1Re-like, the event must sat-
2010 isfy:

- 2011 • Fully-contained and have $dWall > 80\text{cm}$ and $tWall > 170\text{cm}$
- 2012 • Total of one ring which is reconstructed as electron-like with reconstructed

2013 momentum $P_e > 100\text{MeV}$

2014 • Zero decay electrons are associated with the event

2015 • Passes π^0 rejection cut discussed in section 5.2

2016 The zero decay electron cut removes non-CCQE interactions and the π^0
 2017 rejection cut is designed to remove neutral current π^0 background events which
 2018 can be easily reconstructed as 1Re-like events.

2019 The zero decay electron cut removes non-CCQE interactions and the π^0
 2020 rejection cut is designed to remove neutral current π^0 background events which
 2021 can be easily reconstructed as 1Re-like events.

2022 **CC1 π^+ event selection** This event selection is very similar to that of the 1Re
 2023 sample. The only differences are that the dWall and toWall criteria are changed
 2024 to $> 50\text{cm}$ and $> 270\text{cm}$, respectively, and exactly one decay electron is required
 2025 from the π^+ decay.

2026 **1R μ event selection** A 1R μ -like event is determined by the following cuts:

2027 • Fully-contained and have $\text{dWall} > 50\text{cm}$ and $\text{toWall} > 250\text{cm}$

2028 • Total of one ring which is reconstructed as muon-like with reconstructed
 2029 momentum $P_\mu > 200\text{MeV}$

2030 • Fewer than two decay electrons are associated with the event

2031 • Passes π^+ rejection cut discussed in section 5.2

2032 All of these samples are binned in reconstructed neutrino energy. This is
 2033 possible under a particular interaction mode assumption, as the direction from
 2034 the source is known extremely well. For the 1Re-like and 1R μ -like samples,

$$E_\nu^{rec} = \frac{(M_N - V_{nuc})E_l - m_l^2/2 + M_N V_{nuc} - V_{nuc}^2/2 + (M_P^2 + M_N^2)/2}{M_N - V_{nuc} - E_l + P_l \cos(\theta_{beam})}. \quad (6.2)$$

2035 Where M_N , M_P and m_l are the masses of the neutron, proton and outgoing
 2036 lepton, respectively. $V_{nuc} = 27\text{MeV}$ is the binding energy of the oxygen nucleus
 2037 [186], θ_{beam} is the angle between the beam and the direction of the outgoing
 2038 lepton, and E_l and P_l are the energy and momentum of that outgoing lepton.

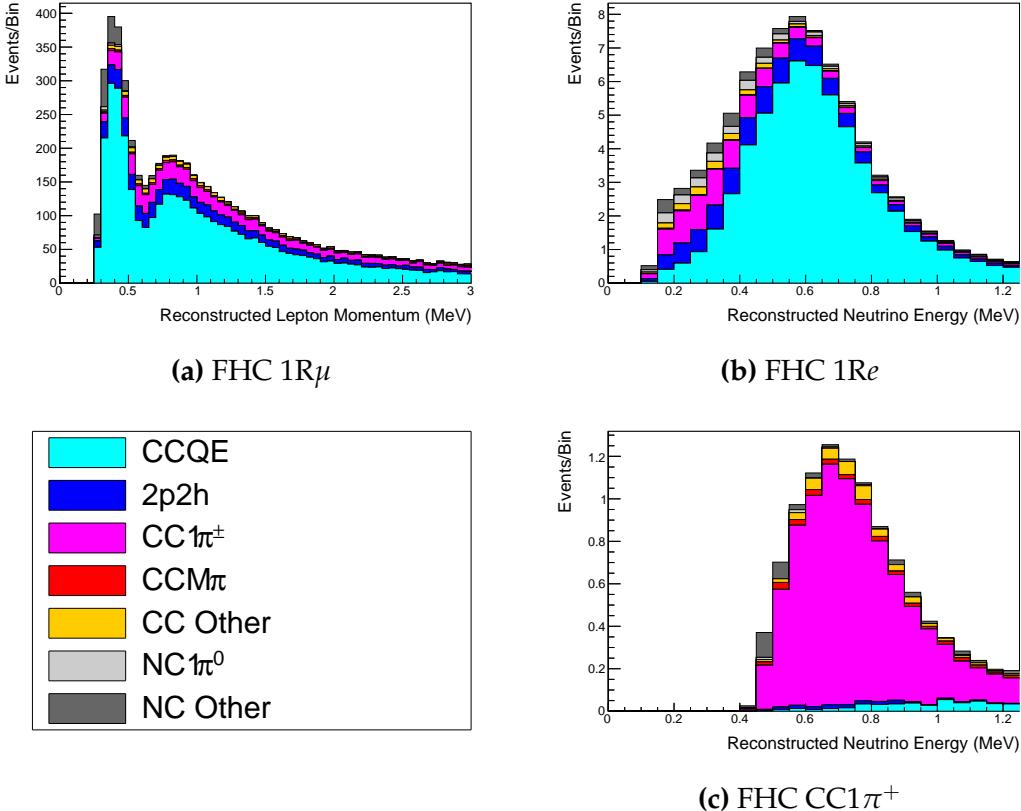


Figure 6.4: The reconstructed neutrino energy, as defined by Equation 6.2 and Equation 6.3, for the 1R μ -like, 1R e -like, and CC1 π^+ -like samples. The AsimovA oscillation parameters are assumed (given in Table 2.2). These samples are the FHC mode samples. For ease of viewing, the 1R μ sample only shows the $0 \leq E_\nu^{rec} < 3.0\text{GeV}$ but the binning extends to 30.0GeV .

2039 The reconstructed neutrino energy of the CC1 π^+ -like events also accounts
 2040 for the delta resonance produced within the interaction,

$$E_\nu^{rec} = \frac{2M_N E_l + M_{\Delta^{++}}^2 - M_N^2 - m_l^2}{2(M_N - E_l + P_l \cos(\theta_{beam}))}. \quad (6.3)$$

2041 Where $M_{\Delta^{++}}$ is the mass of the delta baryon. Binding energy effects are not
 2042 considered as a two-body process, with the delta baryon, is assumed. This follows
 2043 the T2K oscillation analysis presented in [75], although recent developments of

2044 the interaction model in the latest T2K oscillation analysis do include effects
2045 from binding energy in this calculation [200].

2046 The reconstructed neutrino energy for the FHC samples is illustrated in
2047 Figure 6.4. As expected, the $1R\mu$ -like and $1Re$ -like samples are heavily dominated
2048 by CCQE interactions, with smaller contributions from $2p2h$ meson exchange and
2049 resonant pion production interactions. The $CC1\pi^+$ -like sample predominantly
2050 consists of charged current resonant pion production interactions. The $1Re$ -like
2051 and $CC1\pi^+$ -like samples are also binned by the angle between the neutrino beam
2052 and the reconstructed lepton momentum. This is to aid in charged current and
2053 neutral current separation, as indicated in Figure 6.5. This is because the neutral
2054 current backgrounds are predominantly due to π^0 -decays, which decay into two
2055 γ rays. The opening angle of which (alongside the different final state kinematics)
2056 can produce a slightly broader angular distribution compared to the final state
2057 particles originating from charged current ν_e interactions.

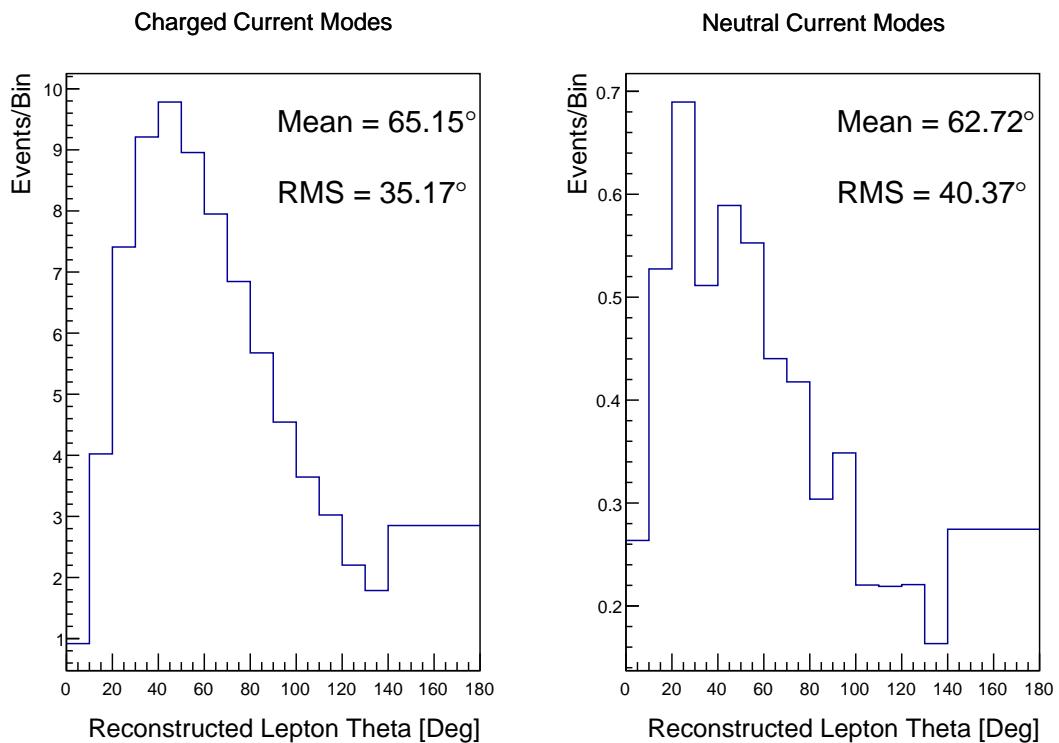


Figure 6.5: The distribution of the angle between the neutrino beam direction and the reconstructed final state lepton, for the FHC 1Re-like sample. The distribution is broken down by neutrino interaction mode into charged current (left) and neutral current (right) components. Asimov A oscillation parameter sets are assumed (given in Table 2.2). The RMS of the charged and neutral current plots are 35.17° and 40.37° , respectively.

2058 6.4 Systematic Uncertainties

2059 The systematic model parameters for this analysis are split into groups, or blocks,
2060 depending on their purpose. They consist of flux uncertainties, neutrino-matter
2061 interaction systematics, and detector efficiencies. There are also uncertainties on
2062 the oscillation parameters to which this analysis is not sensitive, namely Δm_{21}^2
2063 and $\sin^2(\theta_{12})$. These oscillation parameter uncertainties are taken from the 2020
2064 PDG measurements [76]. As described in chapter 4, each model parameter used
2065 within this analysis requires a prior uncertainty. This is provided via separate
2066 covariance matrices for each block. The covariance matrices can include prior
2067 correlations between parameters within a single block, but the separate treatment
2068 means prior correlations can not be included for parameters in different groups.
2069 Some parameters in these models have no reasonably motivated uncertainties
2070 and are assigned flat priors which do not modify the likelihood penalty. In
2071 practice, these flat prior parameters are actually assigned a Gaussian with a
2072 very large width to ensure the covariance matrix is positive definite. They are
2073 then checked at run time to determine if they contribute to the likelihood. The
2074 flux, neutrino interaction, and detector modeling simulations have already been
2075 discussed in section 5.1 and section 5.2. The uncertainties invoked within each
2076 of these models are described below.

2077 6.4.1 Beam Flux

2078 The neutrino beam flux systematics are based upon the uncertainty in the mod-
2079 eling of the components of the beam simulation. This includes the model of
2080 hadron productions and reinteractions, the shape, intensity, and alignment of
2081 the beam with respect to the target, and the uniformity of the magnetic field
2082 produced by the horn, alongside other effects. The uncertainty, as a function
2083 of neutrino energy, is illustrated in Figure 6.6 which includes a depiction of
2084 the total uncertainty as well as the contribution from individual components.
2085 The uncertainty around the peak of the energy distribution ($E_\nu \sim 0.6\text{GeV}$) is

2086 dominated by uncertainties in the beam profile and alignment. Outside of this
2087 region, uncertainties on hadron production dominate the error.

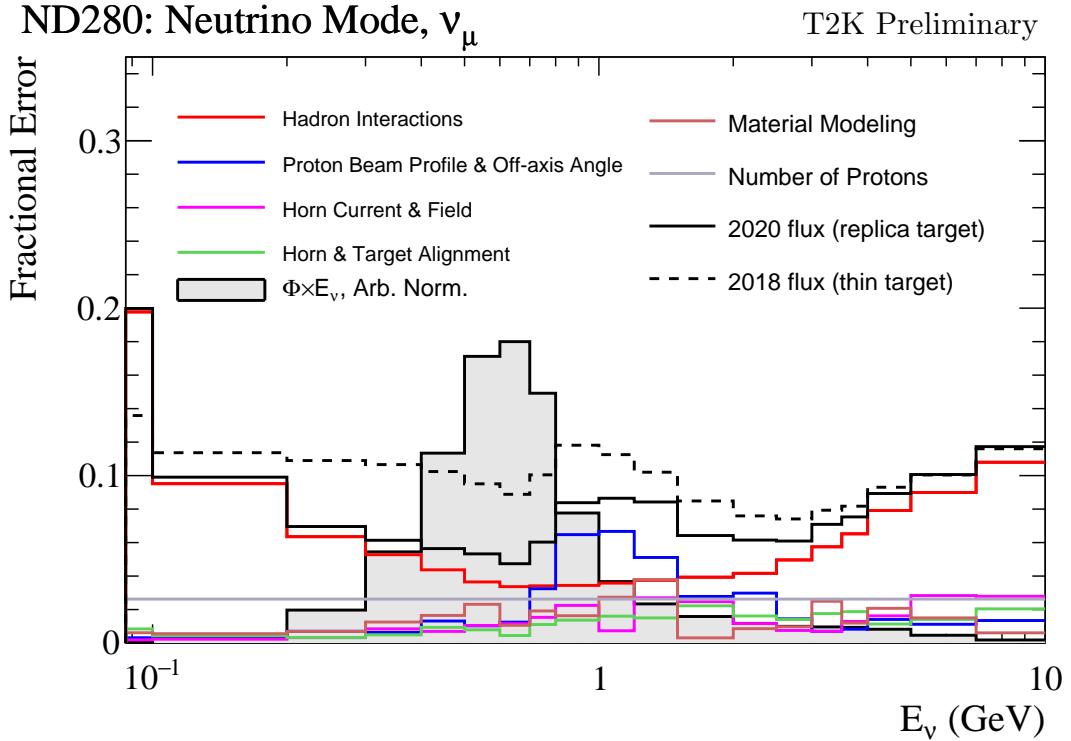


Figure 6.6: The total uncertainty evaluated on the near detector ν_μ flux prediction constrained by the replica-target data, illustrated as a function of neutrino energy. The solid(dashed) line indicates the uncertainty used within this analysis(the T2K 2018 analysis [201]). The solid histogram indicates the neutrino flux as a function of energy. Figure taken from [202].

2088 The beam flux uncertainties are described by one hundred parameters. They
2089 are split between the ND280 and SK detectors and binned by neutrino flavour:
2090 ν_μ , $\bar{\nu}_\mu$, ν_e and $\bar{\nu}_e$. The response is then broken down as a function of neutrino
2091 energy. The bin density in the neutrino energy is the same for the ν_μ in FHC
2092 and $\bar{\nu}_\mu$ in RHC beams, and narrows for neutrino energies close to the oscillation
2093 maximum of $E_\nu = 0.6\text{GeV}$. This binning is specified in Table 6.5. All of these
2094 systematic uncertainties are applied as normalisation parameters with Gaussian
2095 priors centered at 1.0 and error specified from a covariance matrix provided
2096 by the T2K beam group [202].

Neutrino Flavour	Sign	Neutrino Energy Bin Edges (GeV)
μ	Right	0., 0.4, 0.5, 0.6, 0.7, 1., 1.5, 2.5, 3.5, 5., 7., 30.
μ	Wrong	0., 0.7, 1., 1.5, 2.5, 30.
e	Right	0., 0.5, 0.7, 0.8, 1.5, 2.5, 4., 30.
e	Wrong	0., 2.5, 30.

Table 6.5: The neutrino energy binning for the different neutrino flavours. “Right” sign indicates neutrinos in the FHC beam and antineutrinos in the RHC beam. “Wrong” sign indicates antineutrinos in the FHC beam and neutrinos in the RHC beam. The binning of the detector response is identical for the FHC and RHC modes as well as at ND280 and SK.

2097 6.4.2 Atmospheric Flux

2098 The atmospheric neutrino flux is modeled by the HKKM model [51]. 16 systematic
 2099 uncertainties are applied to control the normalisation of each neutrino flavour,
 2100 energy, and direction. They are summarised below:

- 2101 • **Absolute Normalisation:** The overall normalisation of each neutrino flavour
 2102 is controlled by two independent systematic uncertainties, for $E_\nu < 1\text{GeV}$
 2103 and $E_\nu > 1\text{GeV}$, respectively. This is driven mostly by hadronic interaction
 2104 uncertainties for the production of pions and kaons [51]. The strength of
 2105 the response is dependent upon the neutrino energy. The uncertainty is
 2106 parameterized following Figure 11 in [51].
- 2107 • **Relative Normalisation:** Uncertainties on the ratio of $(\nu_\mu + \bar{\nu}_\mu) / (\nu_e + \bar{\nu}_e)$
 2108 are controlled by the difference between the HKKM model [51], FLUKA
 2109 [54] and Bartol models [50]. Three independent parameters are applied in
 2110 the energy ranges: $E_\nu < 1\text{GeV}$, $1\text{GeV} < E_\nu < 10\text{GeV}$, and $E_\nu > 10\text{GeV}$.
- 2111 • **$\nu/\bar{\nu}$ Normalisation:** The uncertainties in the π^+/π^- (and kaon equivalent)
 2112 production uncertainties in the flux of $\nu/\bar{\nu}$. The response is applied using
 2113 the same methodology as the relative normalisation parameters.
- 2114 • **Up/Down and Vertical/Horizontal Ratio:** Similar to the above two sys-
 2115 tematics, the difference between the HKKM, FLUKA, and Bartol model

2116 predictions, as a function of $\cos(\theta_Z)$, is used to control the normalisation of
2117 events as a function of zenith angle.

- 2118 • **K/ π Ratio:** Higher energy neutrinos ($E_\nu > 10\text{GeV}$) mostly originate in
2119 kaon decay. Measurements of the ratio of K/ π production [203] are used to
2120 control the systematic uncertainty of the expected ratio of pion and kaon
2121 production.
- 2122 • **Solar Activity:** As the 11-year solar cycle can affect the Earth's magnetic
2123 field, the flux of primary cosmic rays varies across the same period. The
2124 uncertainty is calculated by taking a ± 1 year variation, equating to a 10%
2125 uncertainty for the SK-IV period.
- 2126 • **Atmospheric Density:** The height of the interaction of the primary cosmic
2127 rays is dependent upon the atmospheric density. The HKKM assumes the
2128 US standard 1976 [153] profile. This systematic controls the uncertainty in
2129 that model.

2130 The total uncertainty is dominated by the absolute and relative normalisation
2131 parameters. The effect of which is illustrated in Figure 6.7. Generally, the
2132 uncertainty is large at low energy, reducing to $O(10\%)$ around the peak of the
2133 flux distribution and then increasing once the neutrino energy exceeds 10GeV.

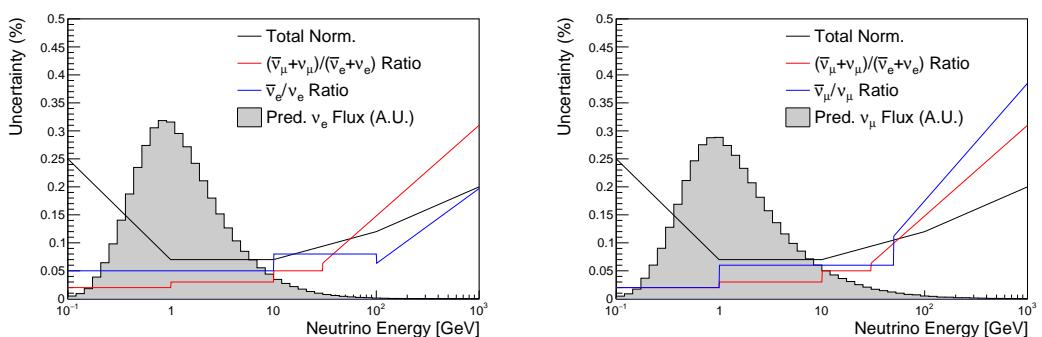


Figure 6.7: The uncertainty evaluated on the atmospheric ν_e (left) and ν_μ (right) flux predictions. The absolute normalisation and flavour ratio uncertainties are given. The solid histogram indicates the neutrino flux as a function of energy.

2134 Updates to the HKKM and Bartol models are underway [158] to use a similar
2135 tuning technique to that used in the beam flux predictions. After those updates,
2136 it may be possible to include correlations in the hadron production uncertainty
2137 systematics for beam and atmospheric flux predictions.

2138 6.4.3 Neutrino Interaction

2139 Neutrino interactions in the detectors are modeled by NEUT. The two indepen-
2140 dent oscillation analyses, T2K-only [204] and the SK-only [60], have developed
2141 separate interaction models. To maximise sensitivity out of this simultaneous
2142 beam and atmospheric analysis, a correlated interaction model has been defined
2143 in [205]. Where applicable, correlations allow the systematic uncertainties applied
2144 to the atmospheric samples to be constrained by near detector neutrino beam
2145 measurements. This can lead to stronger sensitivity to oscillation parameters
2146 as compared to an uncorrelated model.

2147 The low-energy T2K systematic model has a more sophisticated treatment
2148 of CCQE, 2p2h, and CCRES uncertainties, where extensive comparisons of
2149 this model have been performed to external data [204]. However, the model
2150 is not designed for high-energy atmospheric events, like those illustrated in
2151 Figure 5.11. Therefore the high energy systematic model from the SK-only
2152 analysis is implemented for the relevant multi-GeV, PC, and up- μ samples.
2153 The T2K CCQE model is more sophisticated so it has been implemented for
2154 all samples within this analysis, where separate low-energy and high-energy
2155 dials have been implemented. The low-energy dials are constrained by the near
2156 detector measurements and are uncorrelated to their high-energy counterparts.
2157 The author of this thesis was responsible for implementing and validating the
2158 combined cross-section model as documented in [205, 206].

2159 The high energy systematic model includes parameters developed from
2160 comparisons of Nieves and Rein-Seghal models which affect resonant pion
2161 producing interactions, comparisons of the GRV98 and CKMT models which
2162 control DIS interactions, and hadron multiplicity measurements which modulate

the normalisation of multi-pion producing events. The uncertainty on the ν_τ cross-section is particularly large and is controlled by a 25% normalisation uncertainty. These uncertainties are applied via normalisation or shape parameters. The former linearly scales the weight of all affected Monte-Carlo events, whereas the latter can increase or decrease a particular event's weight depending on its neutrino energy and mode of interaction. The response of the shape parameters is defined by third-order polynomial splines which return a weight for a particular neutrino energy. To reduce computational resources for the far detector fit, the response is binned by neutrino energy and sample binning: lepton momentum and cosine zenith binning for atmospheric splined responses and reconstructed neutrino energy and direction binning for beam samples. In total, 17 normalisation and 15 shape parameters are included in the high-energy model within this analysis.

Figure 6.8 indicates the predicted neutrino energy distribution for both beam and subGeV atmospheric samples. There is clearly significant overlap in neutrino energy between the subGeV atmospheric and beam samples, allowing similar kinematics in the final state particles. Figure 6.9 illustrates the fractional contribution of the different interaction modes per sample.

Comparing beam and atmospheric samples which target CCQE interactions (S.G. e-like 0de, S.G. μ -like [0,1]de, [FHC,RHC] 1R μ -like and [FHC,RHC] 1R e-like samples), there is a very similar contribution of CCQE, CC 2p2h, and CC1 π^\pm interactions. The samples which target CC1 π^\pm interactions, (S.G. e-like 0de, S.G. μ -like 2de and FHC 1R+1d.e e-like) also consist of very similar mode interactions.

As a consequence of the similarity in energy and mode contributions, correlating the systematic model between the beam and subGeV atmospheric samples ensures that this analysis attains the largest sensitivity to oscillation parameters while still ensuring neutrino interaction systematics are correctly accounted for. Due to its more sophisticated CCQE and 2p2h model, the T2K systematic model was chosen as the basis of the correlated model.

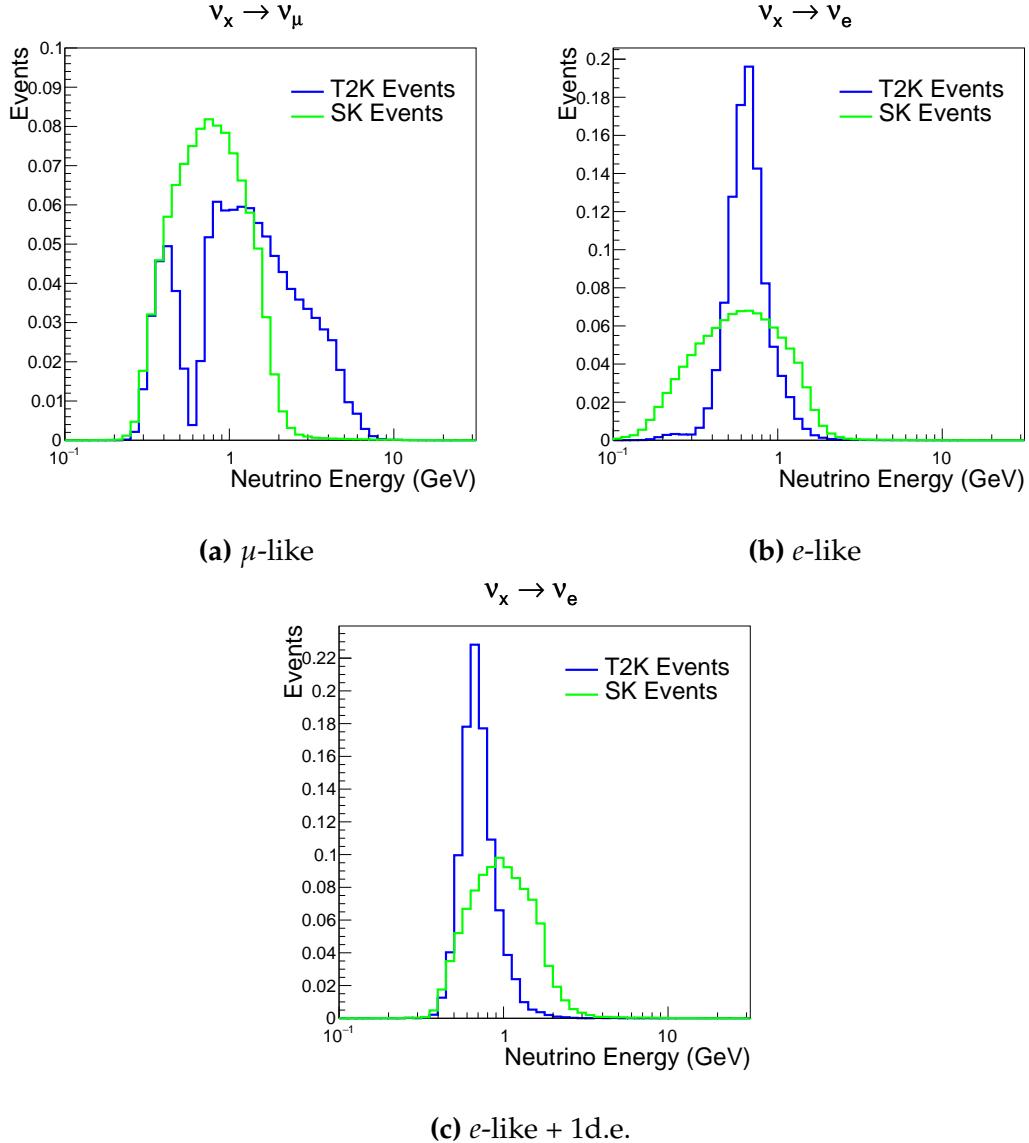


Figure 6.8: The predicted neutrino energy distribution for subGeV atmospheric and beam samples. FHC and RHC beam samples are summed together Asimov A oscillation parameters are assumed (given in Table 2.2). Beam and atmospheric samples with similar cuts are compared against one another.

2193 The T2K systematic model [204] is applied in a similar methodology to the
 2194 SK model parameters. It consists of 19 shape parameters and 24 normalisation
 2195 parameters. Four additional parameters, which model the uncertainty in the
 2196 binding energy, are applied in a way to shift the momentum of the lepton emitted
 2197 from a nucleus. This controls the uncertainty specified on the 27MeV binding
 2198 energy assumed within Equation 6.2. The majority of these parameters are

	CC QE	CC 2p2h	CC $1\pi^\pm$	CC $M\pi$	CC Other	NC $1\pi^0$	NC $1\pi^\pm$	NC $M\pi$	NC Coh.	NC Other
FHC 1R+1d.e. e-like	0.04	0.02	0.83	0.03	0.04	0.01	0.01	0.01	0.00	0.01
RHC 1R e-like	0.62	0.12	0.11	0.01	0.02	0.06	0.01	0.01	0.01	0.04
FHC 1R e-like	0.68	0.12	0.10	0.00	0.02	0.04	0.01	0.00	0.00	0.02
RHC 1R μ -like	0.62	0.13	0.17	0.02	0.03	0.00	0.02	0.00	0.00	0.00
FHC 1R μ -like	0.62	0.12	0.16	0.02	0.03	0.00	0.03	0.00	0.00	0.00
S.G. π^0 -like	0.05	0.01	0.02	0.00	0.01	0.68	0.06	0.07	0.06	0.04
S.G. μ -like 2de	0.04	0.01	0.80	0.10	0.04	0.00	0.00	0.00	0.00	0.00
S.G. μ -like 1de	0.72	0.11	0.12	0.01	0.02	0.00	0.01	0.00	0.00	0.00
S.G. μ -like 0de	0.68	0.11	0.10	0.01	0.02	0.01	0.05	0.01	0.00	0.02
S.G. e-like 1de	0.05	0.01	0.75	0.10	0.05	0.00	0.01	0.02	0.00	0.01
S.G. e-like 0de	0.73	0.11	0.10	0.01	0.02	0.02	0.00	0.00	0.00	0.00

Figure 6.9: The interaction mode contribution of each sample given as a fraction of the total event rate in that sample. Asimov A oscillation parameters are assumed (given in Table 2.2). The Charged Current (CC) modes are broken into quasi-elastic (QE), 2p2h, resonant charged pion production ($1\pi^\pm$), multi-pion production ($M\pi$), and other interaction categories. Neutral Current (NC) interaction modes are given in interaction mode categories: π^0 production, resonant charged pion production, multi-pion production, and others.

2199 assigned a Gaussian prior uncertainty. Those that have no reasonably motivated
 2200 uncertainty, or those which have not been fit to external data, are assigned a
 2201 flat prior which does not affect the penalty term.

2202 On top of the combination of the SK and T2K interaction models, several
 2203 other parameters have been specifically developed for the joint oscillation anal-
 2204 ysis. The majority of the atmospheric samples' δ_{CP} sensitivity comes from the
 2205 normalisation of subGeV electron-like events. These are modeled using a spectral
 2206 function to approximate the nuclear ground state. However, the near detector is
 2207 not able to constrain the model so an additional systematic is introduced which
 2208 models an alternative Continuous Random Phase Approximation (CRPA) nuclear
 2209 ground state. This dial approximates the event weights if a CRPA model had

been assumed rather than a spectral function. This dial only applies to ν_e and $\bar{\nu}_e$ as the near detector does not constraint ν_e cross-section measurements. It is applied as a shape parameter.

Further additions to the model have been introduced due to the inclusion of the subGeV π^0 atmospheric sample. This particularly targets charged current and neutral current π^0 producing interactions to help constrain the systematic uncertainties. Therefore, an uncertainty that affects neutral current resonant π^0 production is incorporated into this analysis. Comparisons of NEUT's NC resonant pion production predictions have been made to MiniBooNE [207] data and a consistent 16% to 21% underprediction is observed [205]. Consequently, a conservative 30% normalisation parameter is invoked.

Down-going events are mostly insensitive to oscillation parameters and can act similar to the near detector within an accelerator experiment (Details will be discussed in chapter 7). This region of phase space can act as a sideband and allows the cross-section model and near detector constraint to be studied. The distribution of events in this region is calculated using the technique outlined in subsection 4.3.4. The results are illustrated in Figure 6.10. For CCQE-targeting samples, the application of the near detector constraint is well within the statistical fluctuation of the down-going data. This means there is no significant tension is observed between the data and the Monte Carlo prediction after the near detector constraint is applied. This is not the case for samples with target CCRES interactions. The electron-like data is consistent with the constrained prediction at high reconstructed momenta but diverges at lower momentum, whereas the muon-like sample is under-predicted throughout the range of momenta. To combat this disagreement, an additional cross-section systematic dial, specifically designed to inflate the low pion momentum systematics was developed in [205]. This is a shape parameter implemented through a splined response.

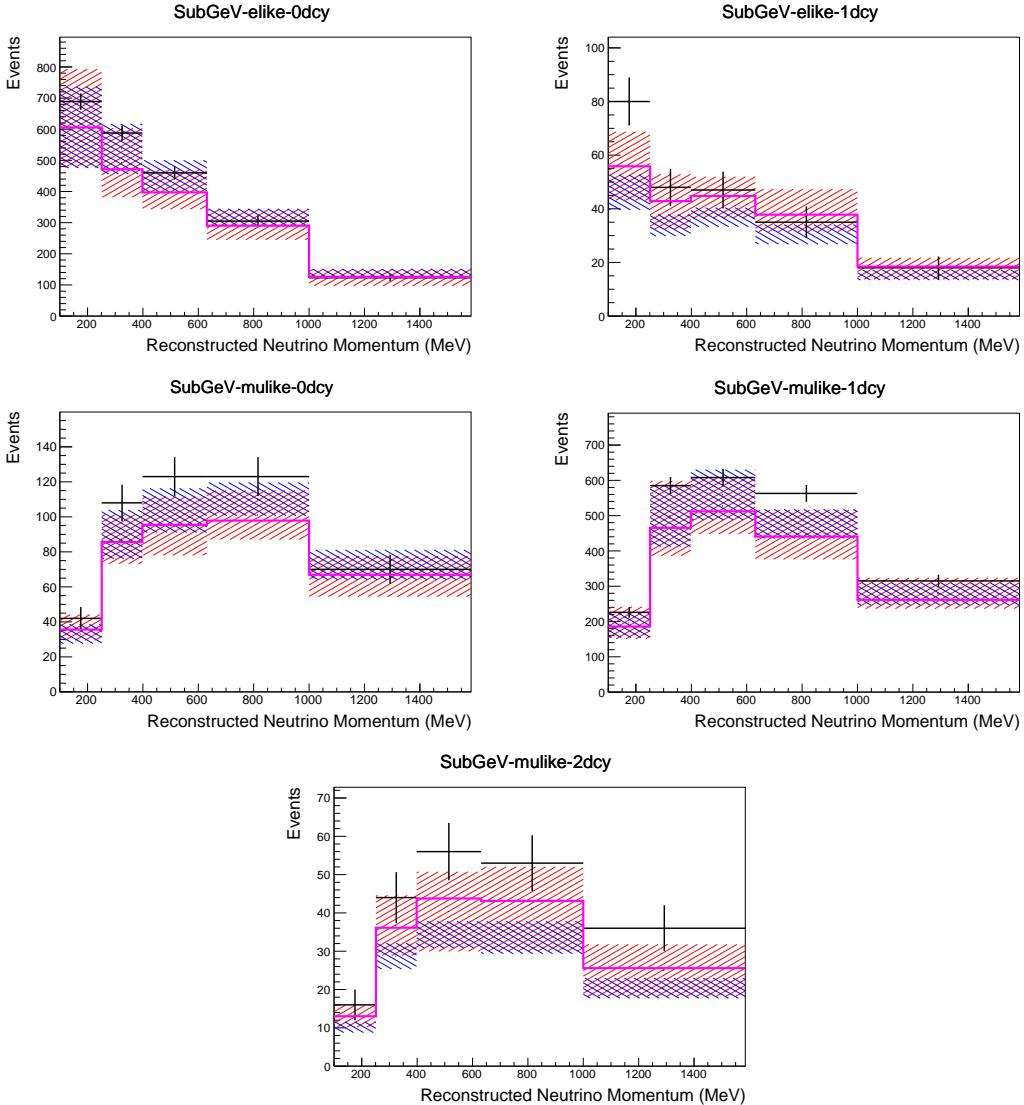


Figure 6.10: Down-going atmospheric subGeV single-ring samples comparing the mean and error of the pre-fit and post-fit Monte Carlo predictions in red and blue, respectively. The magenta histogram illustrates the Monte Carlo prediction using the generated dial values. The black points illustrate the down-going data with statistical errors given. The mean and errors of the Monte Carlo predictions are calculated by the techniques documented in subsection 4.3.4. The pre-fit spectrum is calculated by throwing the cross-section and atmospheric flux dial values from the pre-fit covariance matrix. The post-fit spectrum is calculated by sampling the cross-section dial values from an ND fit MCMC chain, whilst still throwing the atmospheric flux dials from the pre-fit covariance.

2237 6.4.4 Near Detector

2238 The systematics applied due to uncertainties arising from the response of the near
 2239 detector is documented in [132]. The response is described by 574 normalisation
 2240 parameters binned in the selected sample as well as momentum and angle,

P_μ and $\cos(\theta_\mu)$, of the final-state muon. These are applied via a covariance matrix with each parameter being assigned a Gaussian prior from that covariance matrix. These normalisation parameters are built from underlying systematics, e.g. pion secondary interaction systematics, which are randomly thrown and the variation in each $P_\mu \times \cos(\theta_\mu)$ bin is determined. Two thousand throws are evaluated and a covariance matrix response is created. This allows significant correlations between FGD1 and FGD2 samples, as well as adjacent $P_\mu \times \cos(\theta_\mu)$ bins. Statistical uncertainties are accounted for by including fluctuations of each event's weight from a Poisson distribution.

Similar to the cross-section systematics, MaCh3 and BANFF are used to constrain the uncertainty of these systematics through independent validations. Each fitter generates a post-fit covariance matrix which is compared and passed to the far-detector oscillation analysis working group. As the analysis presented within this thesis uses the MaCh3 framework, a joint oscillation analysis fit of all three sets of samples and their respective systematics is performed.

6.4.5 Far Detector

Two configurations of the far detector systematic model implementation have been considered. Firstly, the far detector systematic uncertainties for beam and atmospheric samples are taken from their respective analysis inputs, denoted “official inputs” analysis, with no correlations assumed between the beam and atmospheric samples. The beam- and atmospheric-specific inputs are documented in subsubsection 6.4.5.1 and subsubsection 6.4.5.2. Secondly, an alternative detector model has been developed which correlates the response of the SK detector systematics between the beam and atmospheric samples. Here, the distribution of parameters used for applying event cuts (e.g. electron-muon PID separation) is modified within the fit. It follows a similar methodology to the beam far detector systematics implementation but performs a joint fit of the beam and atmospheric data. This alternative implementation is detailed in subsubsection 6.4.5.3.

2270 **6.4.5.1 Beam Samples**

2271 There are 45 systematics which describe the response of the far detector to
 2272 beam events [186], split into 44 normalisation parameters and one energy scale
 2273 systematic. The energy scale systematic is applied as a multiplicative scaling
 2274 of the reconstructed neutrino energy. It is estimated from data-to-Monte Carlo
 2275 differences in the stopping muon sample in [188] and found to be 2.1%. The
 2276 normalisation parameters are assigned a Gaussian error centered at one with
 2277 width taken from a covariance matrix. A detailed breakdown of the generation
 2278 of the covariance matrix is found in [199]. To build the covariance matrix, a fit
 2279 is performed on atmospheric data which has been selected using beam sample
 2280 selection cuts. These cuts use the variables, L^i , where the index i is detailed in
 2281 Table 6.6. Each L^i is a smear, α , and shift, β parameter such that,

$$L_j^i \rightarrow \bar{L}_j^i = \alpha_j^i L + \beta_j^i. \quad (6.4)$$

2282 Where L_j^i (\bar{L}_j^i) correspond to nominal(varied) PID cut parameters given in
 2283 Table 6.6. The shift and smear parameters are nuisance parameters with no prior
 2284 constraints. They are binned by final-state topology, j , where the binning is given
 2285 in Table 6.7. The final-state topology binning is because the detector will respond
 2286 differently to events that have one or multiple rings. For example, the detector
 2287 will be able to distinguish single-ring events better than two overlapping ring
 2288 events, resulting in different systematic uncertainty for one-ring events compared
 2289 to two-ring events. This approach is used to allow the cut parameter distributions
 2290 to be modified within the fit, allowing for better data to Monte Carlo agreement.

Cut Variable	Parameter Name
0	<code>fitQun e/mu PID</code>
1	<code>fitQun e/pi0 PID</code>
2	<code>fitQun mu/pi PID</code>
3	<code>fitQun Ring-Counting Parameter</code>

Table 6.6: List of cut variables that are included within the shift/smear fit documented in [199].

Category	Description
1e	Only one electron above Cherenkov threshold in the final state
1 μ	Only one muon above Cherenkov threshold in the final state
1e+other	One electron and one or more other charged particles above Cherenkov threshold in the final state
1 μ +other	One muon and one or more other charged particles above Cherenkov threshold in the final state
1 π^0	Only one π^0 in the final state
1 π^\pm or 1p	Only one hadron (typically charged pion or proton) in the final state
Other	Any other final state

Table 6.7: Reconstructed event topology categories on which the SK detector systematics [199] are based.

2291 The mis-modeling of π^0 events is also considered. If one of the two rings
 2292 from a π^0 event is missed, this will be reconstructed as a CC ν_e -like event. This
 2293 is one of the largest systematics hindering the electron neutrino appearance
 2294 analyses. Consequently, additional systematics have been introduced to con-
 2295 strain the mis-modeling of π^0 events in SK, binned by reconstructed neutrino
 2296 energy. To evaluate this systematic uncertainty, a set of “hybrid- π^0 ” samples is
 2297 constructed. These events are built by overlaying one electron-like ring from
 2298 the SK atmospheric neutrino samples or decay electron ring from a stopping
 2299 cosmic ray muon with one simulated photon ring. Both rings are chosen so
 2300 that momenta and opening angle follow the decay kinematics of NC π^0 events
 2301 from the T2K-MC. Hybrid- π^0 Monte Carlo samples with both rings from the
 2302 SK Monte Carlo are produced to compare with the hybrid- π^0 data samples and
 2303 the difference in the fraction of events that pass the ν_e selection criteria is used
 2304 to assign the systematic errors. In order to investigate any data to Monte Carlo
 2305 differences that may originate from either the higher energy ring or lower energy
 2306 ring, two samples are built; a sample in which the electron constitutes the higher
 2307 energy ring from the π^0 decay (called the primary sample) and another one in
 2308 which it constitutes the lower energy ring (called the secondary sample). The
 2309 standard T2K ν_e fitQun event selection criteria are used to select events.

2310 Final contributions to the covariance matrix are determined by supplemen-
 2311 tary uncertainties obtained by comparing stopping muon data to Monte Carlo

prediction, as first introduced in section 5.2. The efficiency of tagging decay electrons is estimated by the stopping muon data to Monte Carlo differences by comparing the number of one decay electron events to the number of events with one or fewer decay electrons. Similarly, the rate at which fake decay electrons are reconstructed by `fiTQun` is estimated by comparing the number of two decay electron events to the number of events with one or two reconstructed decay electrons. The two sources of systematics are added in quadrature weighted by the number of events with one true decay electron yielding a 0.2% systematic uncertainty. A fiducial volume systematic of $\pm 2.5\text{cm}$ which corresponds to a 0.5% shift in the normalisation of events is also applied. Additional normalisation uncertainties based on neutrino flavour and interaction mode are also defined in [186, 208, 209].

Two additional sources of uncertainty are included: secondary and photoneuclear interactions. These are estimated by varying the underlying parameters are building a distribution of sample event rates. These contributions are then added in quadrature to the above covariance matrix. The final uncertainty on the SK detector systematics are provided in Figure 6.11.

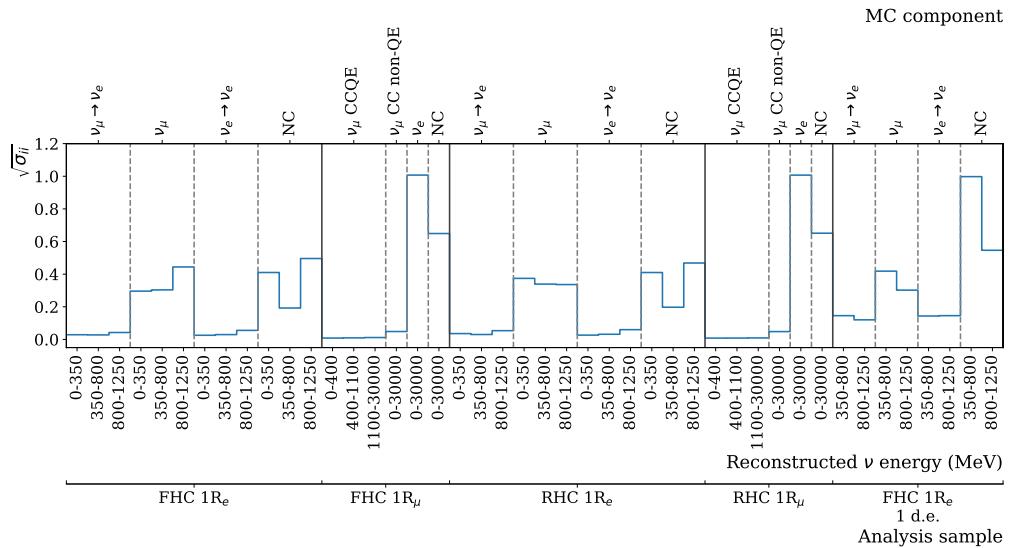


Figure 6.11: The fractional uncertainty on each of the 44 parameters describing the SK detector systematics (The energy scale systematic is neglected). The parameters are split by sample, oscillation channel, interaction mode and reconstructed neutrino energy.

6.4.5.2 Atmospheric Samples

The detector systematics for atmospheric samples, documented in [88], are split into two sub-groups: those which are related to particle identification and ring counting systematics, and those which are related to calibration, separation, and reduction uncertainties.

The particle identification systematics consist of five parameters. The ring separation systematic enforces an anti-correlated response between the single-ring and multi-ring samples. This is implemented as a fractional increase/decrease in the overall normalisation of each sample, depending on the distance to the nearest wall from an event's vertex. The coefficients of the normalisation are estimated prior to the fit and depend on the particular atmospheric sample. Two electron-muon separation systematics are included within this model which anti-correlates the response of the electron-like and muon-like samples: one for single-ring events and another for multi-ring events.

The multi-ring electron-like separation likelihood, discussed in section 6.1, encodes the ability of the detector to separate neutrino from anti-neutrino events. Two normalisation parameters vary the relative normalisation of multi-ring ν_e and $\bar{\nu}_e$ samples whilst keeping a consistent overall event rate.

There are 22 systematics related to calibration measurements, including effects from backgrounds, reduction, and showering effects. They are documented in [88] and are briefly summarised in Table 6.8. They are applied via normalisation parameters, with the separation systematics requiring the conservation of event rate across all samples.

6.4.5.3 Correlated Detector Model

A complete uncertainty model of the SK detector would be able to determine the systematic shift on the sample spectra for a variation of the underlying parameters, e.g. PMT angular acceptance. However, this is computationally intensive, requiring Monte Carlo predictions to be made for each plausible variation. Consequently, an effective parameter model has been utilised for

Index	Description
0	Partially contained reduction
1	Fully contained reduction
2	Separation of fully contained and partially contained events
3	Separation of stopping and through-going partially contained events in top of detector
4	Separation of stopping and through-going partially contained events in barrel of detector
5	Separation of stopping and through-going partially contained events in bottom of detector
6	Background due to cosmic rays
7	Background due to flasher events
8	Vertex systematic moving events into and out of fiducial volume
9	Upward going muon event reduction
10	Separation of stopping and through-going in upward going muon events
11	Energy systematic in upward going muon events
12	Reconstruction of the path length of upward going muon events
13	Separation of showering and non-showering upward going muon events
14	Background of stopping upward going muon events
15	Background of non-showering through-going upward going muon events
16	Background of showering through-going upward going muon events
17	Efficiency of tagging two rings from π^0 decay
18	Efficiency of decay electron tagging
19	Background from downgoing cosmic muons
20	Asymmetry of energy deposition in tank
21	Energy scale deposition

Table 6.8: Sources of systematic errors specified within the grouped into the “calibration” systematics model.

2358 a correlated detector model following from the T2K-only model implementation
 2359 documented in subsubsection 6.4.5.1. It correlates the detector systematics
 2360 between the far-detector beam and subGeV atmospheric samples due to their
 2361 similar energies and interaction types. As there are no equivalent beam samples,
 2362 the multi-GeV, multiring, PC, and Up- μ samples will be subject to the particle
 2363 identification systematics implementation as described in subsubsection 6.4.5.2
 2364 rather than using this correlated detector model. The calibration systematics also
 2365 described in the aforementioned chapter still apply to all atmospheric samples.
 2366 The correlated detector model utilises the same smear and shift parameters
 2367 documented in subsubsection 6.4.5.1, split by final state topology. Beyond this,

2368 the shift and smear parameters are split by visible energy deposited within the
 2369 detector, with binning specified in Table 6.9. This is because atmospheric events
 2370 are categorised by subGeV and multi-GeV events based on visible energy, so
 2371 this splitting is required when correlating the systematic model for beam and
 2372 atmospheric events. Alongside the technical requirement, higher energy events
 2373 will be better reconstructed due to fractionally less noise within the detector. As
 2374 a result of the inclusion of visible energy binning, Equation 6.4 becomes

$$L_{jk}^i \rightarrow \bar{L}_{jk}^i = \alpha_{jk}^i L + \beta_{jk}^i, \quad (6.5)$$

2375 where k is the visible energy bin.

Index	Range (MeV)
0	$30 \geq E_{vis} > 300$
1	$300 \geq E_{vis} > 700$
2	$700 \geq E_{vis} > 1330$
3	$E_{vis} \geq 1330$

Table 6.9: Visible energy binning for which the correlated SK detector systematics are based

2376 The implementation of this systematic model takes the events reconstructed
 2377 values of the cut parameters, modifies them by the particular shift and smear
 2378 parameter for that event, and then re-applies event selection. This causes event
 2379 migration, which is a new feature incorporated into the MaCh3 framework which
 2380 is only achievable due to the event-by-event reweighting scheme.

2381 Particular care has to be taken when varying the ring counting parameter.
 2382 This is because the number of rings is a finite value (one-ring, two-ring, etc.)
 2383 which can not be continuously varied through this shift and smear technique.
 2384 Consequently a continuous ring counting parameter, RC_i , is calculated for the
 2385 i^{th} event, following the definition in [185]: the preferred likelihoods from all
 2386 considered one-ring (L_{1R}) and two-ring (L_{2R}) fits are determined. The difference

2387 is computed as $\Delta_{LLH} = \log(L_{1R}) - \log(L_{2R})$. The ring counting parameter is
 2388 then defined as

$$RC_i = \text{sgn}(\Delta_{LLH}) \times \sqrt{|\Delta_{LLH}|}, \quad (6.6)$$

2389 where $\text{sgn}(x) = x/|x|$. This ring counting parameter corresponds to an
 2390 intermediate likelihood value used within the `fitQun` algorithm to decide the
 2391 number of rings associated with a particular event. However, fake-ring merging
 2392 algorithms are applied after this likelihood value is used. Consequently, this
 2393 ring counting parameter does not always exactly correspond to the number of
 2394 reconstructed rings. This can be seen in Figure 6.12.

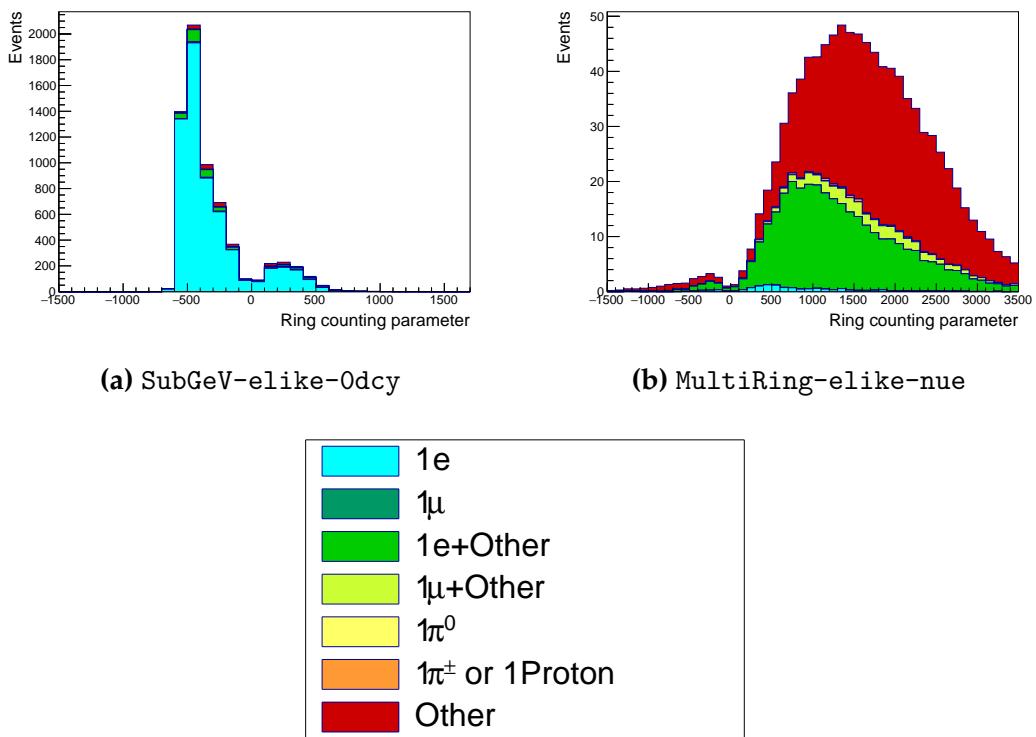


Figure 6.12: The ring counting parameter as defined in Equation 6.6 for the SubGeV-elike-0dcy and MultiRing-elike-nue samples.

2395 As the `fitQun` algorithm does not provide a likelihood value after the fake-
 2396 ring algorithms have been applied, the ring counting parameter distribution is
 2397 correlated to the final number of reconstructed rings through “maps”. These

2398 are two-dimensional distributions of the ring counting parameter and the final
 2399 number of reconstructed rings. An example is illustrated in Figure 6.13. In
 2400 principle, the `fitQun` reconstruction algorithm should be re-run after the variation
 2401 in the ring counting parameter. However, this is not computationally viable.
 2402 Therefore the “maps” are used as a reweighting template.

2403 The maps are split by final state topology and true neutrino flavour and
 2404 all `fitQun`-reconstructed Monte Carlo events are used to fill them. The maps
 2405 are row-normalised to represent the probability of X rings for a given RC_i
 2406 value. Prior to the oscillation fit, an event’s nominal weight is calculated as
 2407 $W^i(N_{Rings}^i, L_{jk}^i)$, where N_{Rings}^i is the reconstructed number of rings for the i^{th}
 2408 event and $W^i(x, y)$ is the bin content in map associated with the i^{th} event, where
 2409 x number of rings and y is ring counting parameter. Then during the fit, the
 2410 value of $R = W^i(N_{Rings}^i, \bar{L}_{jk}^i) / W^i(N_{Rings}^i, L_{jk}^i)$ is calculated as the event weight
 2411 for the i^{th} event. This is the only cut variable that uses a reweighting technique
 2412 rather than event migration.

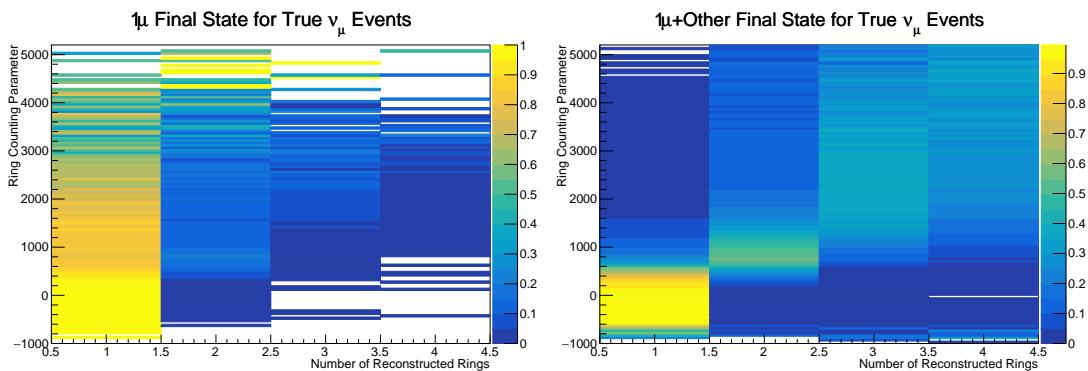


Figure 6.13: The ring counting parameter, defined in Equation 6.6, as a function of the number of reconstructed rings as found by the `fitQun` reconstruction algorithm. Left: true ν_μ events with only one muon above the Cherenkov threshold in the final state. Right: true ν_μ events with one muon and at least one other charged particle above the Cherenkov threshold in the final state.

2413 The π^0 systematics introduced in subsection 6.4.4 are applied via a covariance
 2414 matrix. This is not possible in the alternative model as no covariance matrix
 2415 is used. Thus, the implementation of the π^0 systematics has been modified.
 2416 The inputs from the hybrid π^0 sample are included via the use of “ χ^2 maps”,

which are two-dimensional histograms in α_{jk}^i and β_{jk}^i parameters over some range. Illustrative examples of the χ^2 maps are given in Figure 6.14. Due to their nature, the shift and smear parameters are typically very correlated. A map is produced for each cut parameter given in Table 6.6 and for each visible energy bin given in Table 6.9.

The maps are filled through the χ^2 comparison of the hybrid π^0 Monte Carlo and data in the particle identification parameters documented in Table 6.6. The Monte Carlo distribution is modified by the α_{jk}^i and β_{jk}^i scaling, whilst cross-section and flux nuisance parameters are thrown from their prior uncertainties. The χ^2 between the scaled Monte Carlo and data is calculated and the relevant point in the χ^2 map is filled.

The implementation within this alternative detector model is to add the bin contents of the maps, for the relevant values of the α_{jk}^i and β_{jk}^i parameters, to the likelihood penalty. Only $1\pi^0$ final state topology shift and smear parameters use this prior uncertainty.

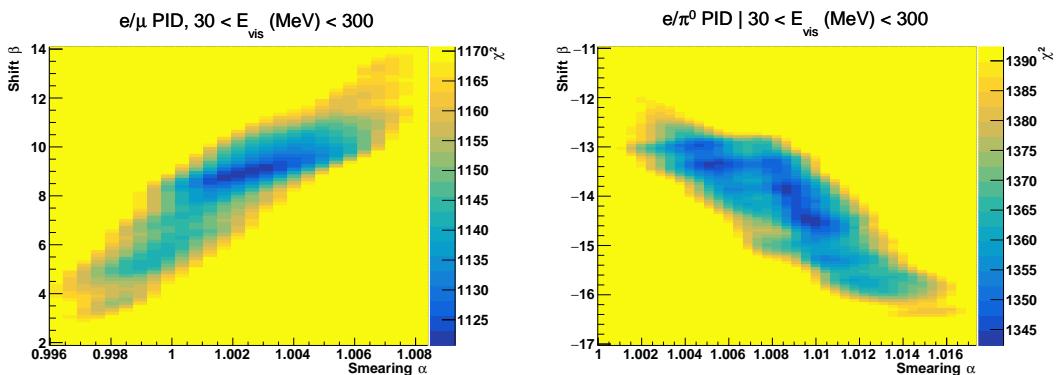


Figure 6.14: The χ^2 between the hybrid- π^0 Monte Carlo and data samples, as a function of smear (α) and shift (β) parameters, for events which have $1\pi^0$ final state topology. Left: Electron-muon separation PID parameter for events with $30 \leq E_{\text{vis}}(\text{MeV}) < 300$. Right: Electron- π^0 separation PID parameter for events with $30 \leq E_{\text{vis}}(\text{MeV}) < 300$.

Similarly, the implementation of the supplementary systematics documented in subsubsection 6.4.5.1 needs to be modified. A new framework [210] was built in tandem between the author of this thesis and the T2K-SK working group [186] so the additional parameters can be incorporated into the MaCh3 framework. These are applied as normalisation parameters, depending on the particular

2437 interaction mode, number of tagged decay electrons, and whether the primary
2438 particle generated Cherenkov light. They are assigned Gaussian uncertainties
2439 with widths described by a covariance matrix. Furthermore, the secondary
2440 interaction and photo-nuclear effects need to be accounted for in this detector
2441 model using a different implementation than that in subsubsection 6.4.5.1. This
2442 was done by including a shape parameter for each of the secondary interactions
2443 and the photo-nuclear systematic parameters.

2444 There are a total of 224 α_{jk}^i and β_{jk}^i parameters, of which 32 have prior
2445 constraints from the hybrid π^0 samples.

2446 One final complexity of this correlated detector model is that the two sets
2447 of samples, beam and subGeV atmospheric, use slightly different parameters
2448 to distinguish electron and muon-like events. The T2K samples use the value
2449 of $\log(L_e/L_\mu)$ whereas the atmospheric samples use the value of $\log(L_e/L_\pi)$,
2450 where L_X is the likelihood for hypothesis X. This is because the T2K fits use
2451 single-ring f iTQun fitting techniques, whereas multi-ring fits are applied to the
2452 atmospheric samples where only the electron and pion hypothesis are considered.
2453 The correlation between the two likelihood ratios is illustrated in Figure 6.15. As
2454 discussed in section 5.2, the pion hypothesis is a very good approximation of the
2455 muon hypothesis due to their similar mass. Consequently, using the same shift
2456 and smear parameters correlated between the beam and subGeV atmospheric
2457 samples is deemed a good approximation.

2458 6.5 Likelihood Calculation

2459 This analysis performs a joint oscillation parameter fit of the ND280 beam
2460 samples, the T2K far detector beam samples, and the SK atmospheric samples
2461 introduced in this chapter.

2462 Once the Monte Carlo predictions of each beam and atmospheric sample
2463 have been built, a likelihood needs to be constructed. This is done by comparing
2464 the binned Monte Carlo prediction to binned data. The Monte Carlo prediction
2465 is calculated at a particular point, $\vec{\theta}$, in the model parameter space such that

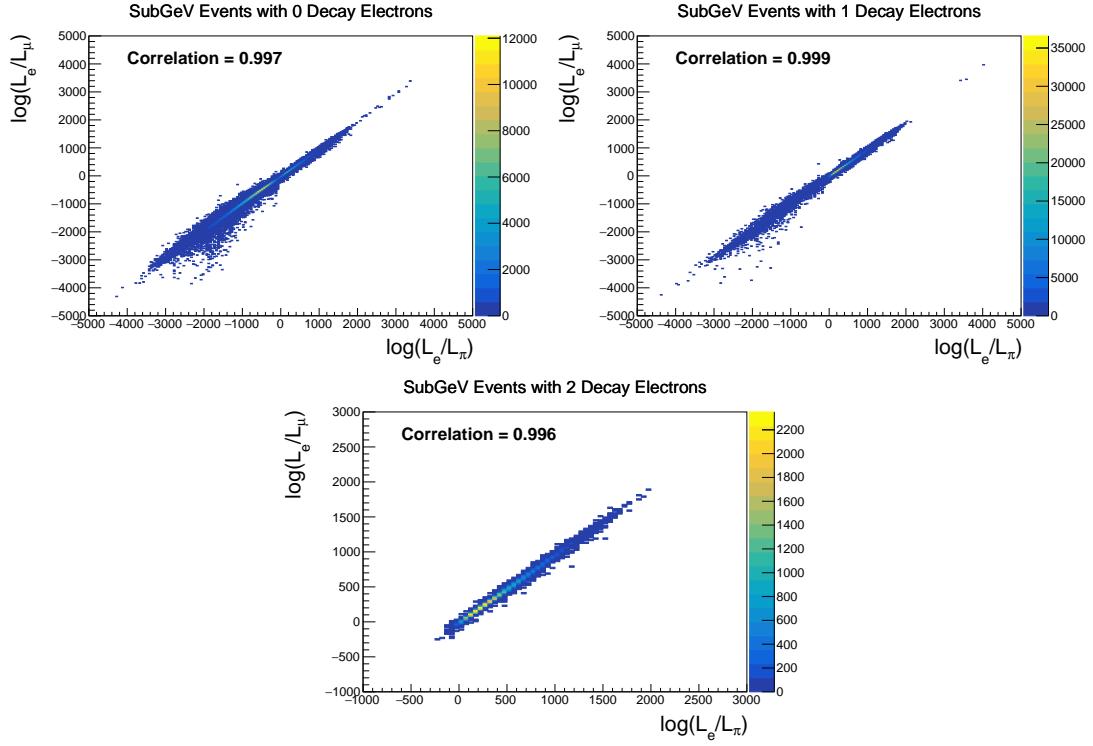


Figure 6.15: The distribution of $\log(L_e/L_\mu)$ compared to $\log(L_e/L_\pi)$ for subGeV events with zero (top left), one (top right) or two (bottom) decay electrons. The correlation in the distribution is calculated as 0.997, 0.999 and 0.996, respectively.

2466 $N_i^{MC} = N_i^{MC}(\vec{\theta})$, where N_i represents the bin content of the i^{th} bin. The data
2467 and Monte Carlo spectra are represented by N_i^D and N_i^{MC} , respectively. The bin
2468 contents for the beam near detector, beam far detector and atmospheric samples
2469 are denoted with ND , FD , and Atm , respectively. The binning index, i , runs
2470 over all the bins within a sample. Taking the FHC1Rmu far detector sample as
2471 an example, the binning index runs over all the reconstructed neutrino energy
2472 bins. The likelihood calculation between the data and the Monte Carlo prediction
2473 for a particular bin follows a Poisson distribution, where the data is treated as
2474 a fluctuation of the simulation.

2475 The data can consist of either real data or an ‘Asimov’ Monte Carlo prediction,
2476 which is typically used for sensitivity studies and denoted ‘Asimov data’. The
2477 process for building Asimov data is as follows. The Monte Carlo prediction is
2478 reweighted using a particular set of oscillation parameters (potentially those
2479 listed in Table 2.2) and systematic parameter tune. The resulting spectra for each

sample is then defined to be the Asimov data for that sample. Whilst this results in unphysical non-integer data predictions, it eliminates statistical fluctuations from the data. Therefore, the results of a fit to Asimov data should not include any biases from statistical fluctuations. Furthermore, these results should produce posterior probability distributions consistent with the parameters which were used to make the data prediction. That is to say, the fit results should return the known parameters. Any biases seen would be attributed to correlations between each oscillation parameter and correlations between oscillation and systematic parameters. Consequently, Asimov fit results present the maximum precision at which the oscillation parameters could be measured to.

Following the T2K analysis presented in [75], the likelihood contribution for the near detector samples also includes a Monte Carlo statistical uncertainty term, derived from the Barlow and Beeston statistical treatment [211, 212]. It includes a contribution to the likelihood that which treats the generated Monte Carlo prediction as a statistical fluctuation of the actual true simulation assuming an infinite amount of statistics had been created. The technical implementation of this additional likelihood term is documented in [194] and briefly summarised as follows. The term is defined as,

$$\frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2}, \quad (6.7)$$

where β_i represents a scaling parameter for the i^{th} bin that relates the bin content for the amount of Monte Carlo actually generated N_i^{MC} to the bin content if an infinite amount of Monte Carlo statistics had been generated $N_{i,true}^{MC}$, such that $N_{i,true}^{MC} = \beta_i \times N_i^{MC}$. In the case where a sufficient amount of Monte Carlo statistics had been generated, $\beta_i = 1$. An analytical solution for β_i is given in [194]. Additionally, $\sigma_{\beta_i} = \sqrt{\sum_i w_i^2 / N_i^{MC}}$ where $\sqrt{\sum_i w_i^2}$ represents the sum of the square of the weights of the Monte Carlo events which fall into bin i .

An additional contribution to the likelihood comes from the variation of the systematic model parameters. For those parameters with well-motivated uncertainty estimates, a covariance matrix, V , describes the prior knowledge of

2508 each parameter as well as any correlations between the parameters. Due to a
 2509 technical implementation, a single covariance matrix describes each “block” of
 2510 model parameters, e.g. beam flux systematics. The covariance matrix associated
 2511 with the k^{th} block is denoted V^k . This substitution results in $\vec{\theta} = \sum_k^{N_b} \vec{\theta}^k$ and
 2512 $V = \sum_k^{N_b} V^k$ where N_b denotes the number of blocks. A single covariance
 2513 matrix is provided for: the oscillation parameters, the beam flux parameters,
 2514 the atmospheric flux parameters, the neutrino interaction systematics, the near
 2515 detector parameters, the beam far detector systematics, and the atmospheric far
 2516 detector systematics. The number of parameters in the k^{th} block is defined as $n(k)$.

2517 The equation for the likelihood \mathcal{L} includes all the terms discussed above.

2518 It is defined as,

$$\begin{aligned}
 -\ln(\mathcal{L}) = & \\
 & \sum_i^{\text{NDbins}} N_i^{\text{ND},MC}(\vec{\theta}) - N_i^{\text{ND},D} + N_i^{\text{ND},D} \times \ln \left[N_i^{\text{ND},D} / N_i^{\text{ND},MC}(\vec{\theta}) \right] + \frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2} \\
 & + \sum_i^{\text{FDbins}} N_i^{\text{FD},MC}(\vec{\theta}) - N_i^{\text{FD},D} + N_i^{\text{FD},D} \times \ln \left[N_i^{\text{FD},D} / N_i^{\text{FD},MC}(\vec{\theta}) \right] \\
 & + \sum_i^{\text{Atmbins}} N_i^{\text{Atm},MC}(\vec{\theta}) - N_i^{\text{Atm},D} + N_i^{\text{Atm},D} \times \ln \left[N_i^{\text{Atm},D} / N_i^{\text{Atm},MC}(\vec{\theta}) \right] \\
 & + \frac{1}{2} \sum_k^{N_b} \sum_i^{n(k)} \sum_j^{n(k)} (\vec{\theta}^k)_i (V^k)_{ij}^{-1} (\vec{\theta}^k)_j.
 \end{aligned} \tag{6.8}$$

2519 The negative log-likelihood value is determined at each step of the MCMC
 2520 to build the posterior distribution defined in chapter 4. This value is minimised
 2521 when the Monte Carlo prediction tends towards the data spectrum.

7

2522

2523

Oscillation Probability Calculation

2524 It is important to understand how and where the sensitivity to the oscillation parameters comes from for both atmospheric and beam samples. An
2525 overview of how these samples respond to changes in δ_{CP} , Δm_{32}^2 , and $\sin^2(\theta_{23})$
2526 is given in section 2.5. This section also explains the additional complexities
2527 involved when performing an atmospheric neutrino analysis as compared to
2528 a beam-only analysis.

2529 Without additional techniques, atmospheric sub-GeV upward-going neutrinos ($E_\nu < 1.33\text{GeV}, \cos(\theta_Z) < 0.$) can artificially inflate the sensitivity to $\delta_{CP,zaza}$
2530 due to the quickly varying oscillation probability in this region. Therefore, a
2531 “sub-sampling” approach has been developed to reduce these biases ensuring
2532 accurate and reliable sensitivity measurements. This technique ensures that small-
2533 scale unresolvable features of the oscillation probability have been averaged over
2534 whilst the large-scale features in the oscillation probability are unaffected. The
2535 documentation and validation of this technique are found in section 7.1. The
2536 oscillation probability calculation is computationally intensive due to the large
2537 number of matrix multiplications needed. Consequently, the CUDAProb3 imple-
2538 mentation choice made within the fitting framework, as detailed in section 7.2,
2539 ensures that the analysis can be done in a timely manner.

Whilst the beam neutrinos are assumed to propagate through a constant density slab of material, the density variations through the Earth result in more complex oscillation patterns for atmospheric neutrinos. Furthermore, the uncertainty in the electron density can modify the oscillation probability for the denser core layers of the Earth. The model of the Earth used within this analysis is detailed in section 7.3. This includes information about the official SK-only methodology as well as improvements that have been made to remove some of the approximations used in that analysis. Another complexity of atmospheric neutrino oscillation studies is that the height of production in the atmosphere is not known on an event-by-event basis. An analytical averaging technique that approximates the uncertainty of the oscillation probability has been followed, with the author of this thesis being responsible for the implementation and validation. This implementation of an external technique is described in section 7.4.

7.1 Treatment of Fast Oscillations

As shown in Figure 7.1, atmospheric neutrino oscillations have a significantly more complex structure for upgoing neutrinos with energy below 1GeV. This is because the L/E dependence of the oscillation probability in this region induces rapid variations for small changes in L or E . As discussed in section 2.5, this is also the region in which atmospheric neutrinos have sensitivity to δ_{CP} . In practice, the direction of the neutrino is inferred from the direction of the final state particles traveling in the detector. The correlation between these two directions can be particularly weak for low-energy neutrino interactions. This creates a distinct difference from the beam neutrinos where the position of the source is very precisely known.

As a consequence of the unresolvable structure, an event rate consistent with the averaged oscillation probability is observed in the subGeV upgoing region. This creates a computational problem: A significantly large amount of Monte Carlo statistics would be required to accurately predict the number of events if Monte Carlo averaging was the only technique used. This section describes

2571 the ‘sub-sampling’ approach developed for this analysis and compares it to the
2572 methodology used within the SK-only analysis.

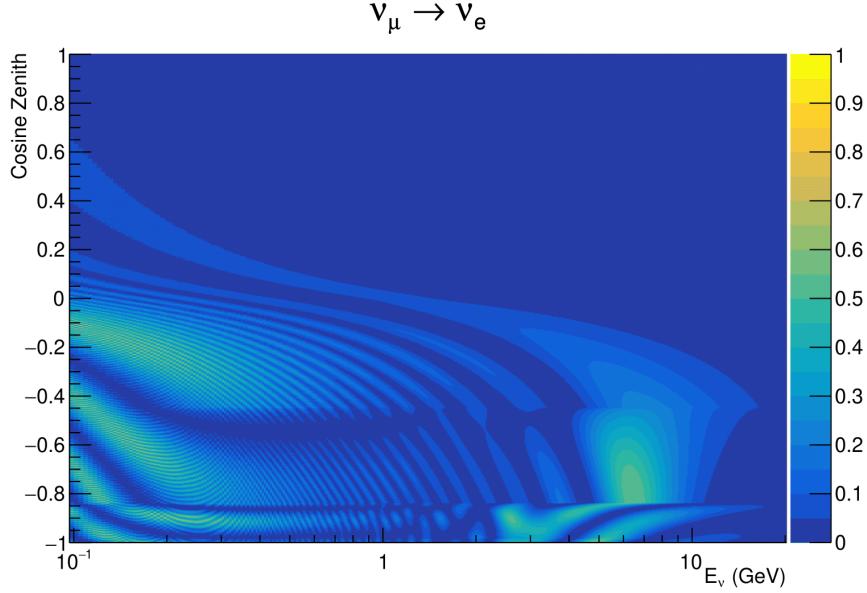


Figure 7.1: The oscillation probability $P(\nu_\mu \rightarrow \nu_e)$, given as a function of neutrino energy and zenith angle, which highlights an example of the “fast” oscillations in the sub-GeV upgoing region.

2573 The official SK-only analysis uses the osc3++ oscillation parameter fitter
2574 [79]. To perform the fast oscillation averaging, it uses a ‘nearest-neighbour’
2575 technique. For a given Monte Carlo neutrino event, the nearest twenty Monte
2576 Carlo neighbours in reconstructed lepton momentum and zenith angle are
2577 found and a distribution of their neutrino energies is built. The RMS, σ , of
2578 this distribution is then used to compute an average oscillation probability for
2579 the given neutrino Monte Carlo event.

2580 For the i^{th} event, the oscillation weight is calculated as

$$W_i = \frac{1}{5}P(E_i, \bar{L}_i) + \frac{1}{5}\sum_{\beta=-1, -0.5, 0.5, 1} P(E_i + \beta\sigma_i, L_\beta), \quad (7.1)$$

2581 where $P(E, L)$ is the oscillation probability calculation for neutrino energy E
2582 and path length L and the two path lengths, \bar{L}_i and L_β are described below. All
2583 of the oscillation probability calculations are performed with a fixed zenith angle
2584 such that the same density profile is used. The uncertainty in the production

height is controlled by using an “average” production height, \bar{L}_i , which represents the average path length computed using twenty production heights taken from the Honda flux model’s prediction [53]. These inputs are provided in 5% intervals of the cumulative distribution function. The value of \bar{L}_i is calculated as:

$$\bar{L}_i = \frac{1}{20} \sum_{j=1}^{20} \sqrt{(R_E + h_j)^2 - R_E^2 (1 - \cos^2 \theta_i)} - R_E \cos \theta_i. \quad (7.2)$$

Where R_E is the Earth’s radius and θ_i is the zenith angle of the i^{th} event. The production heights h_j represent the $(j \times 5)^{th}$ percentile of the cumulative distribution function. L_β values (where the values of β are given in Equation 7.1) are similarly calculated but instead use different combinations of four production heights,

$$\begin{aligned} L_{-1.0} &= \frac{1}{4} L(45, 50, 55, 60), \\ L_{-0.5} &= \frac{1}{4} L(35, 40, 65, 70), \\ L_{+0.5} &= \frac{1}{4} L(25, 30, 75, 68), \\ L_{+1.0} &= \frac{1}{4} L(15, 20, 85, 89). \end{aligned} \quad (7.3)$$

Where $L(i, j, k, l)$ represents the sum of the path lengths with fixed zenith angle and production heights corresponding to the i^{th} , j^{th} , k^{th} and l^{th} percentile of the cumulative distribution function. The values that are taken as β (and values for L_β) are chosen to smooth the oscillation contours in Δm_{32}^2 without incurring loss of sensitivity [213].

This averaging technique works because of the inference between the zenith angle and the reconstructed direction of final state particles in the detector. For low-energy neutrinos, where the resolution of the true neutrino direction is poor, σ_i will be large, resulting in significant averaging effects. Contrary to this, the inferred direction of high-energy neutrinos will be much closer to the true value, meaning that σ_i will be smaller, culminating in small averaging effects.

In practice, these calculations are performed prior to the fit as only oscillation parameters at fixed points are considered. The MCMC technique used in this thesis requires oscillation probabilities to be evaluated at arbitrary parameter values, not known *a priori*. Calculating the five oscillation probabilities per event required by the SK technique is computationally infeasible, so a different averaging technique is used. However, the concept of the averaging technique can be taken from it.

To perform a similar averaging as the SK analysis, a sub-sampling approach using binned oscillograms has been devised. A coarsely binned oscillogram is defined in $\cos(\theta_Z)$ and E_ν . For a given set of oscillation parameters, a single oscillation probability will be assigned to each coarse bin. This value will then apply to all Monte Carlo events which fall into that bin. To assign these oscillation probabilities, the probability is calculated at $N \times N$ points on a grid within a particular bin. This ensemble of oscillation probabilities is averaged to define the coarse bin's oscillation probability, assuming a flat prior in E_ν and $\cos(\theta_Z)$ within the bin. Figure 7.2 illustrates the $N = 2$ example where the assigned value to a coarse bin is the average of the four fine bins which fall in that coarse bin. Whilst the coarse bin edges do not have to be linear on either axis, the sub-division of the fine bins is linear within the range of a coarse bin.

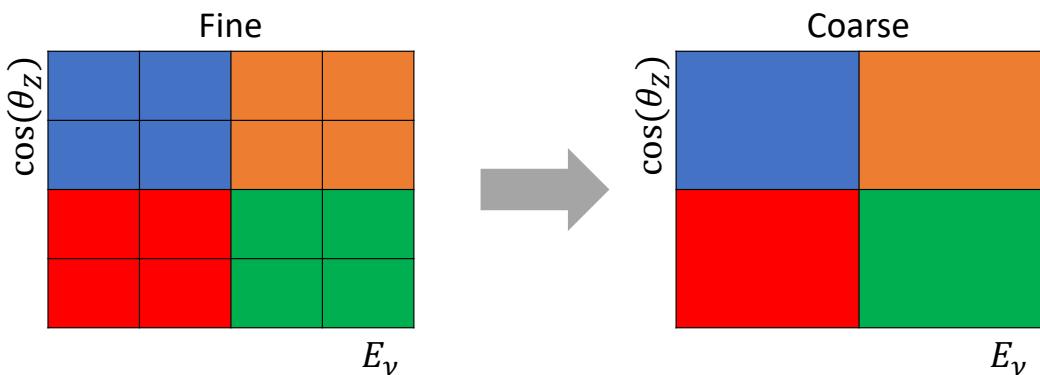


Figure 7.2: Illustration of the averaging procedure for $N = 2$. The oscillation probabilities calculated on the finer left binning are averaged to obtain the oscillation probabilities in the coarser right binning. These averaged oscillation probabilities with the coarser binning are then applied to each event during the fit.

2624 The coarse binning is defined with 67×52 bins in true neutrino energy \times
 2625 cosine zenith. It is picked to be identical to that provided in [213]. In general, the
 2626 binning is logarithmically spaced in neutrino energy but has some hand-picked
 2627 bin edges around the matter resonance to smoothly increased the bin density.
 2628 This is to avoid smearing this region which can be well sampled by the Monte
 2629 Carlo. The cosine zenith binning is approximately linearly spaced across the
 2630 allowable range but the values of layer transitions are hit precisely: -0.8376 (core-
 2631 mantle) and -0.4464 (mantle/transition zone). Bins are spread further apart for
 2632 downgoing events as this is a region unaffected by the fast oscillation wavelengths
 2633 and reduces the total number of calculations required to perform the calculation.

2634 The choice of N is justified based on two studies. Firstly, the variation of event
 2635 rates of each sample is studied as a function of N . For a given set of oscillation
 2636 parameters thrown from the PDG prior constraints (detailed in Table 2.1), the
 2637 oscillation probabilities are calculated using a given value of N . Each sample
 2638 is re-weighted and the event rate is stored. The value of N is scanned from
 2639 1, which corresponds to no averaging, to 19, which corresponds to the largest
 2640 computationally viable subdivision binning. The event rate of each sample at
 2641 large N is expected to converge to a stationary value due to the fine binning fully
 2642 sampling the small-scale structure. Figure 7.3 illustrates this behaviour for the
 2643 SubGeV_elike_0dcy sample for 9 different throws of the oscillation parameters.

2644 Denoting the event rate for one sample for a given throw t at each N by λ_t^N ,
 2645 the average over all considered N values ($\bar{\lambda}_t = \frac{1}{24} \sum_{N=1}^{24} \lambda_t^N$) is computed. The
 2646 variance in the event rate at each N is then calculated as

$$\text{Var}[\lambda^N] = \frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t)^2 - \left[\frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t) \right]^2. \quad (7.4)$$

2647 In practice, the following procedure is undertaken. For a particular throw,
 2648 the difference between the event rate at a particular choice of N and the mean
 2649 of the distribution is calculated. This is illustrated in Figure 7.4. This value
 2650 is then calculated for all the 2000 throws, generating a distribution of $\lambda_t^N - \bar{\lambda}_t$.

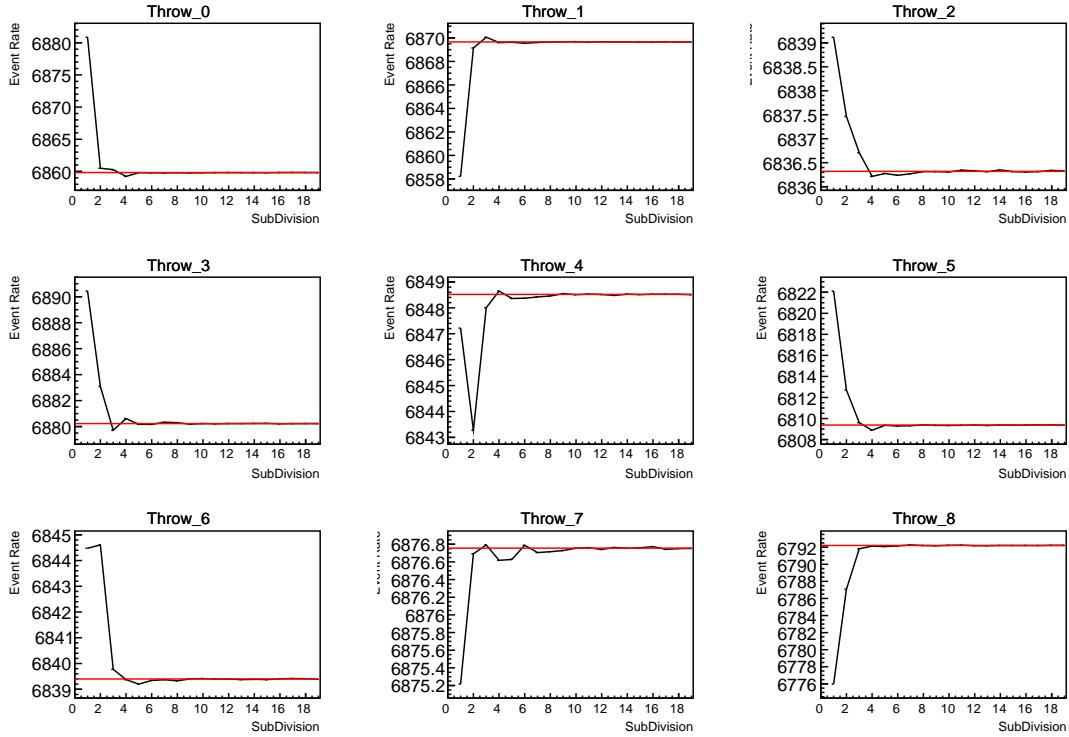


Figure 7.3: Event rate of the SubGeV_elike_0dcy sample as a function of the number of sub-divisions per coarse bin. Each subplot represents the event rate of the sample at a different oscillation parameter set thrown from the PDG priors detailed in Table 2.1. The red line in each subplot represents the mean of the event rate over the different values of sub-divisions for that particular oscillation parameter throw.

2651 This is repeated for each of the values of N considered within this study. The
 2652 distributions of this value, for $N = \{1, 5\}$, are given in Figure 7.5. As expected,
 2653 the distribution gets narrower and tends towards zero for the higher values of N .

2654 The aim of the study is to find the lowest value of N such that this variance
 2655 is below 0.001. This utilises the width of the distributions given in Figure 7.5.
 2656 This is the typical threshold used by T2K fitters to validate systematic imple-
 2657 mentation so has been set as the same criteria. The results of this study for
 2658 each atmospheric sample used within this thesis are illustrated in Figure 7.6 for
 2659 2000 throws of the oscillation parameters. As can be seen, the variance is below
 2660 the threshold at $N = 10$, and is driven primarily by the SubGeV_mulike_1dcy
 2661 and SubGeV_elike_0dcy samples.

2662 The second study to determine the value of N is as follows. The likelihood
 2663 for each sample is computed against an Asimov data set created with Asimov A

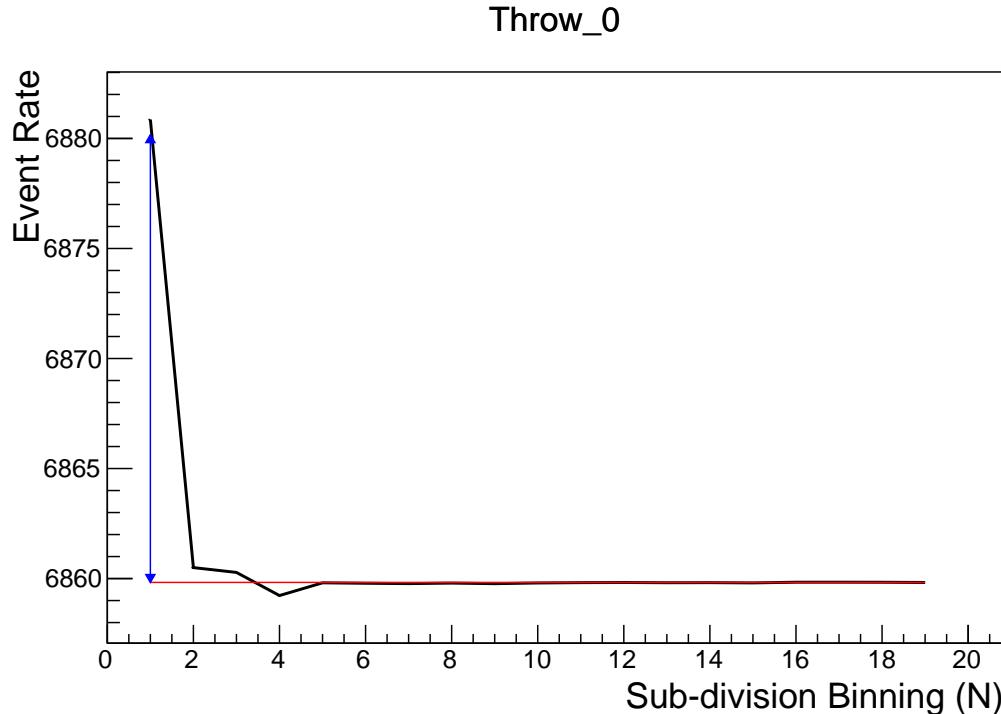


Figure 7.4: Event rate of the SubGeV_elike_0dcy sample, for a particular oscillation parameter throw, as a function of the number of sub-divisions, N , per coarse bin. The difference between the mean event rate (red), $\bar{\lambda}$, and the event rate at $N = 1$, $\lambda^{N=1}$ is defined as $\lambda^N - \bar{\lambda}$ and illustrated by the blue arrow.

2664 oscillation parameters (Table 2.2). Following Equation 7.4, the variance of the log-
 2665 likelihood over all considered N is computed. The results are shown in Figure 7.7.

2666 A choice of $N = 10$ sub-divisions per coarse bin has a variance in both
 2667 event rate and log-likelihood residuals less than the required threshold of 0.001.
 2668 The largest value of the likelihood variance is of order 10^{-7} , corresponding to
 2669 an error on the log-likelihood of about 3×10^{-4} which is small enough to be
 2670 negligible for the oscillation analysis.

2671 Figure 7.8 illustrates the effect of the smearing using $N = 10$. The fast oscilla-
 2672 tions in the sub-GeV upgoing region have been replaced with a normalisation
 2673 effect whilst the large matter resonance structure remains.

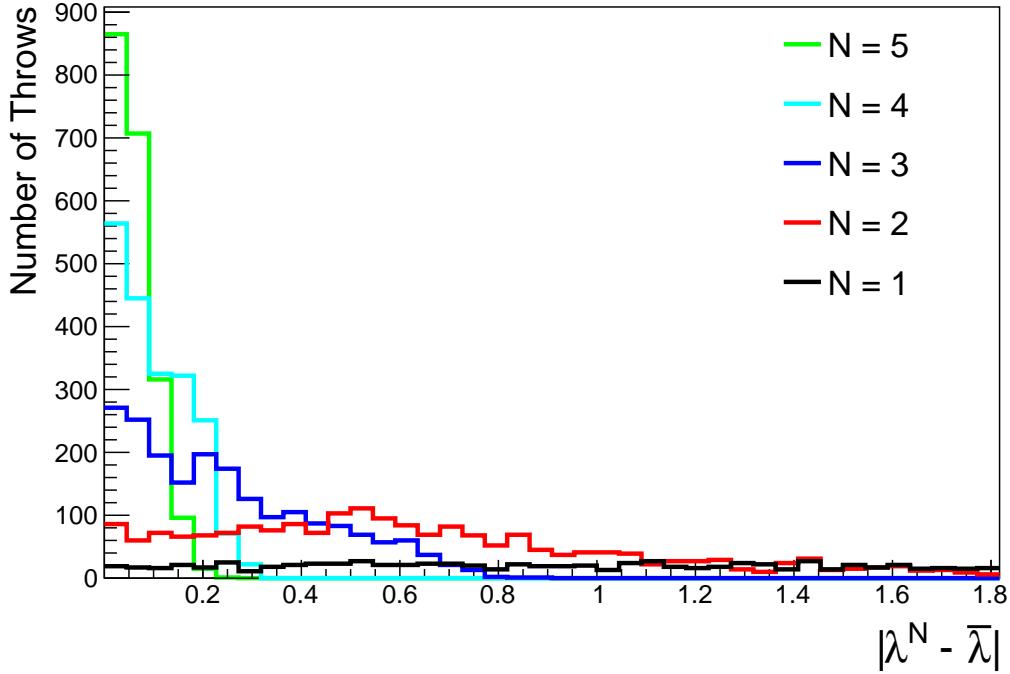


Figure 7.5: The distribution of $\lambda^N - \bar{\lambda}$ for various values of N . As expected, the distribution gets narrower for larger values of N .

2674 7.2 Calculation Engine

2675 As previously discussed in section 7.1, the calculation of oscillation probabilities
 2676 is performed at run-time. Consequently, the time per calculation is crucial for fit
 2677 performance. The initial fitting framework used for this analysis was developed
 2678 with ProbGPU [214]. This is a GPU-only implementation of the prob3 engine
 2679 [215]. It is primarily designed for neutrino propagation in a beam experiment
 2680 (single layer of constant density) with the atmospheric propagation code not
 2681 being used prior to the analysis in this thesis.

2682 Another engine, CUDAProb3 [216], has been interfaced with the fitting frame-
 2683 work used in this analysis. This interfacing was done by the author of this
 2684 thesis. It has been specifically optimised for atmospheric neutrino oscillation
 2685 calculation so does not contain the code to replace the beam oscillation calculation.
 2686 The engine utilises object-orientated techniques as compared to the functional
 2687 implementation of ProbGPU. This allows the energy and cosine zenith arrays to

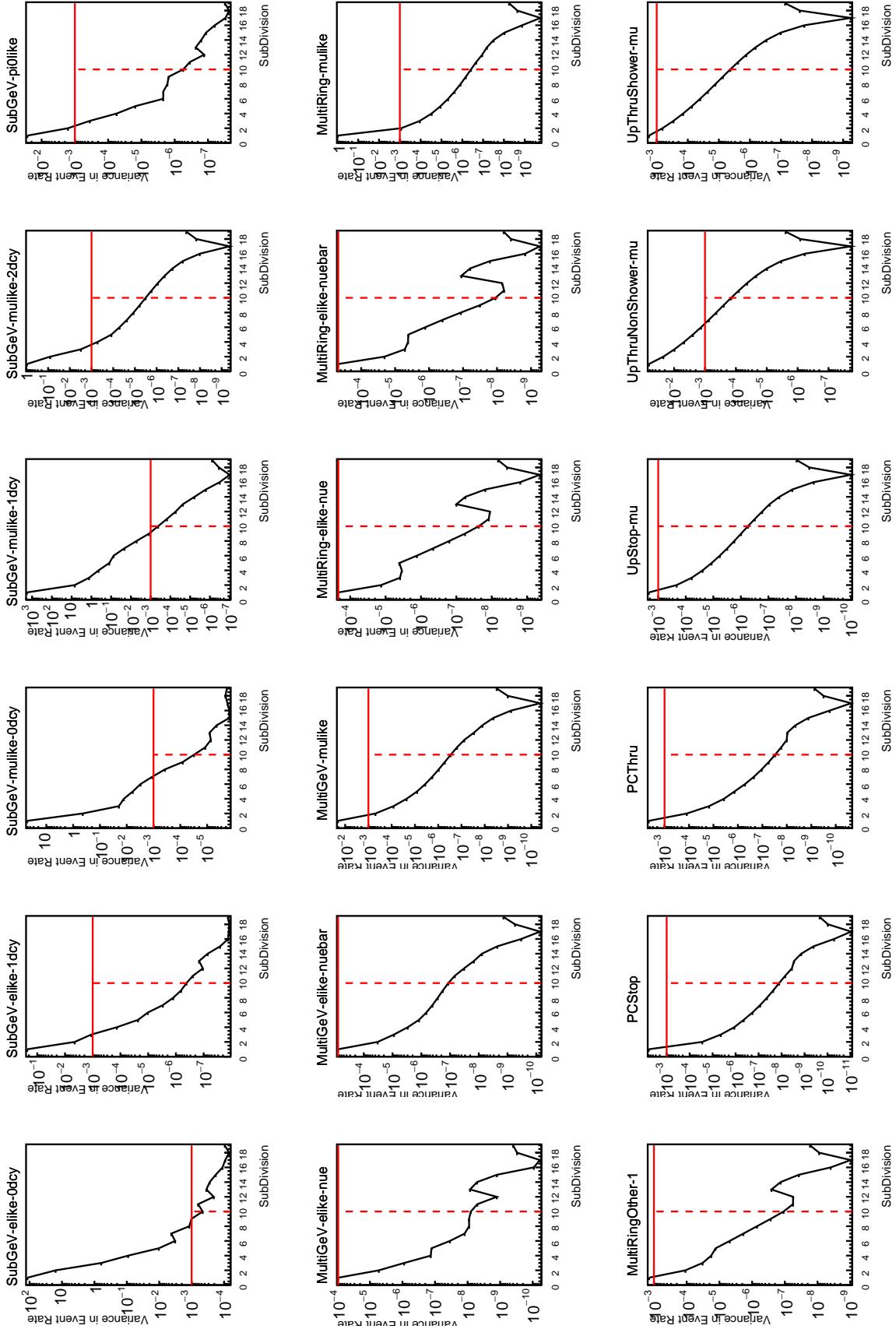


Figure 7.6: Variance of event rate for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line indicates the variance at a sub-division $N = 10$.

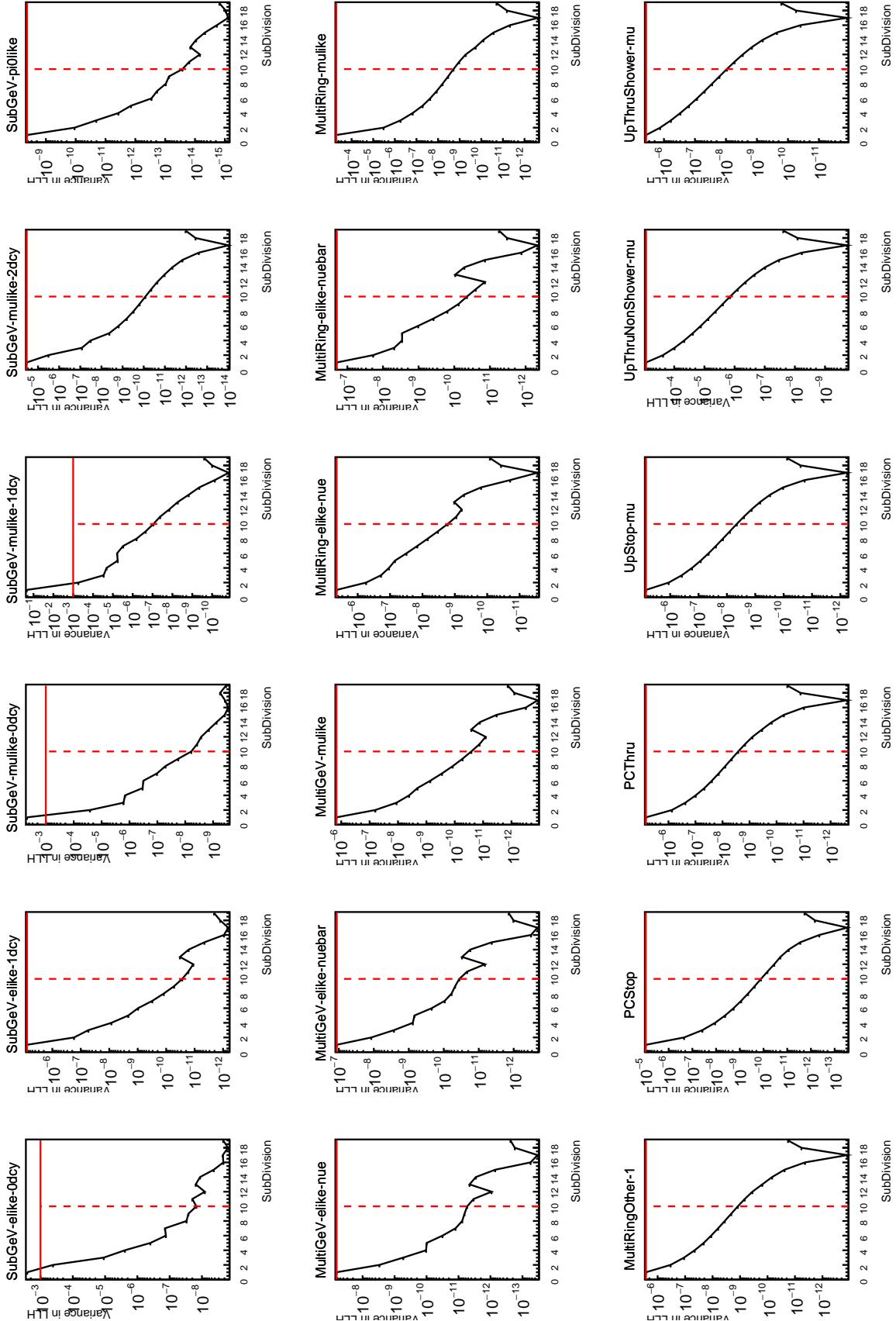


Figure 7.7: Variance of sample likelihood, when compared to ‘Asimov data’ set at Asimov A, for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line is a graphical indication of the variance at a sub-division $N = 10$.

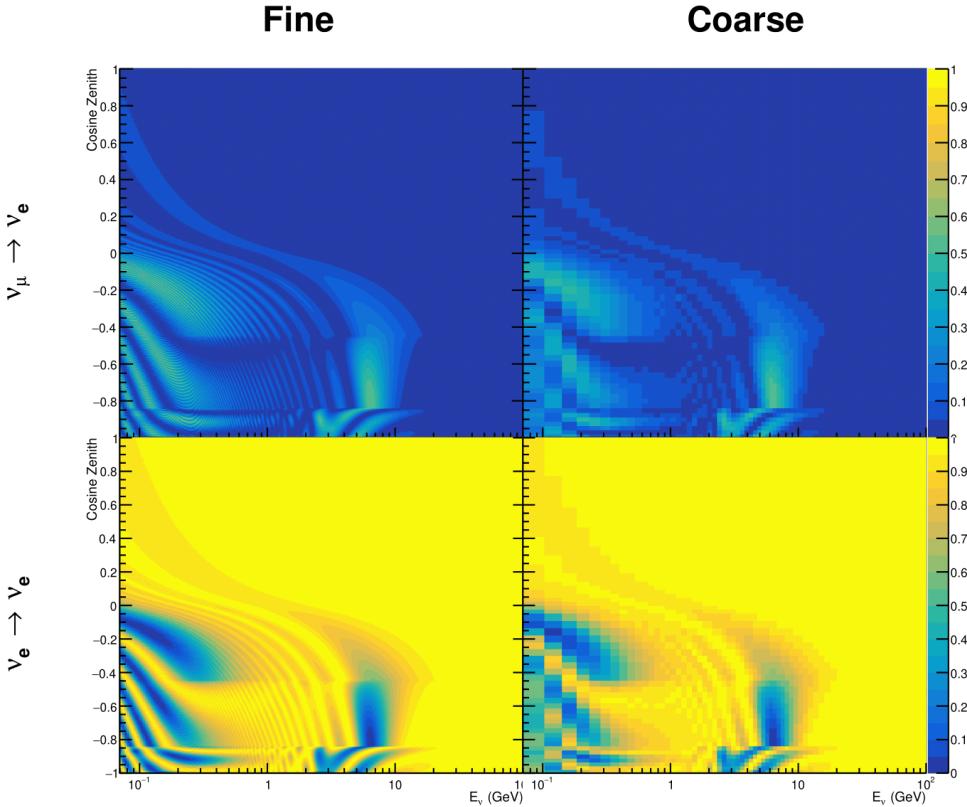


Figure 7.8: The oscillation probability, $P(\nu_\mu \rightarrow \nu_e)$ (top row) and $P(\nu_e \rightarrow \nu_e)$ (bottom row), given as a function of neutrino energy and zenith angle. The left column gives the “fine” binning used to calculate the oscillation probabilities and the right column illustrates the “coarse” binning used to reweight the Monte Carlo events. The fine binning choice is given with $N = 10$, which was determined to be below the threshold from Figure 7.6 and Figure 7.7.

2688 be kept on GPU memory, rather than having to load these arrays onto GPU
 2689 memory for each calculation. Reducing the memory transfer between CPU and
 2690 GPU significantly reduces the time required for calculation. This can be seen
 2691 in Figure 7.9, where the GPU implementation of CUDAProb3 is approximately
 2692 three times faster than the ProbGPU engine.

2693 Another significant advantage of CUDAProb3 is that it contains a CPU multi-
 2694 threaded implementation which is not possible with the ProbGPU or prob3 engines.
 2695 This eliminates the requirement for GPU resources when submitting jobs to batch
 2696 systems. As illustrated in Figure 7.9, the calculation speed depends on the number

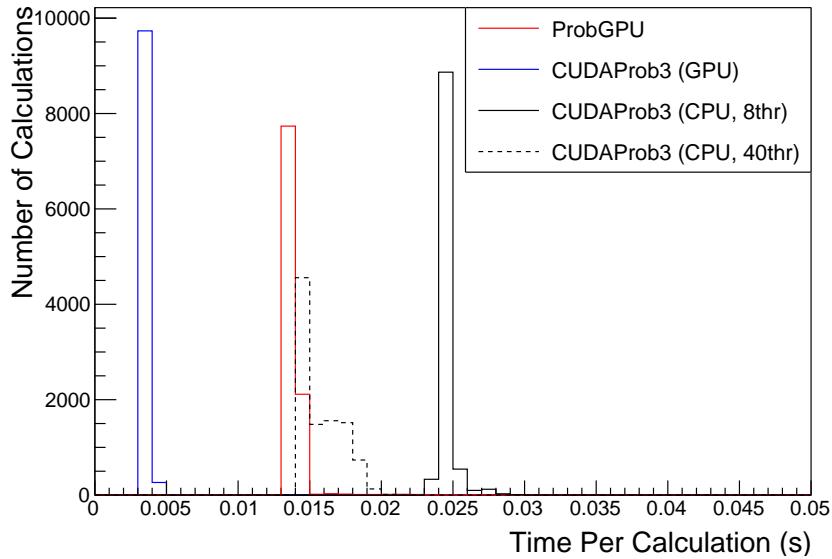


Figure 7.9: The calculation time taken to both calculate the oscillation probabilities and fill the “coarse” oscillograms, following the technique given in section 7.1, for the CUDAProb3 and ProbGPU (Red) calculation engines. CUDAProb3 has both a GPU (Blue) and CPU (Black) implementation, where the CPU implementation is multi-threaded. Therefore, 8-threads (solid) and 40-threads (dashed) configurations have been tested. Prob3, which is a CPU single-thread implementation has a mean step time of 1.142s.

of available threads. Using 8 threads (which is typical of the batch systems being used) is approximately twice as slow as the ProbGPU engine implementation, but would allow the fitting framework to be run on many more resources. This fact is utilised for any SK-only fits but GPU resources are required for any fits which include beam samples due to the ProbGPU requirement. Based on the benefits shown by the implementation in this section, efforts are being placed into including linear propagation for beam neutrino propagation into the engine [217].

7.3 Matter Density Profile

For an experiment observing neutrinos propagating through the Earth, a model of the Earth’s density profile is required. The model used within this analysis is based on the Preliminary Reference Earth Model (PREM) [80], as illustrated in Figure 2.8. Table 2.3 documents the density and radii of the layers used within the constant density approximation used by the SK-only analysis [79]. The

density measurements provided in the PREM model are provided in terms of mass density, whereas neutrino oscillations are sensitive to the electron number density. This value can be computed as the product of the chemical composition, or the Z/A value, and the mass density of each layer. Currently, the only way to measure the chemical composition value for layers close to the Earth's core is through neutrino oscillations. The chemical composition of the upper layers of the Earth's Mantle and the Transition zone is well known due to it being predominantly pyrolite which has a chemical composition value of 0.496 [218]. The chemical composition dial for the core layers is set to a value of 0.468, as calculated in [219]. As this value is less well known, it is assigned a Gaussian error with a standard deviation equivalent to the difference in chemical composition in core and mantle layers. Figure 7.10 illustrates the effect of moving from the $Z/A = 0.5$ method which is used in the official SK-only analysis to these more precise values.

The beam oscillation probability in this thesis uses a baseline of 295km, density 2.6g/cm^3 , and chemical composition 0.5 as is done by the official T2K-only analysis [220].

For a neutrino with given $E_\nu, \cos(\theta_Z)$, the oscillation probability calculation engine must be passed a list of the matter regions that the neutrino traversed, with the path length and fixed density in each region. However, a neutrino passing through the earth experiences a range of radii, and thus a range of densities, in each region. In the SK-only analysis, the earth density model used is piecewise-constant, thereby ignoring this effect. For this thesis, the density values for the calculation engine are found by averaging the earth density along the neutrino's path in each layer,

$$\langle \rho \rangle_i = \frac{1}{t_{i+1} - t_i} \int_{t_i}^{t_{i+1}} \rho(t) dt, \quad (7.5)$$

where t_i are the intersection points between each layer and t is the path length of the trajectory across the layer. This leads to an improved approximation. For this averaging, the simplification of the PREM model developed in [221] is

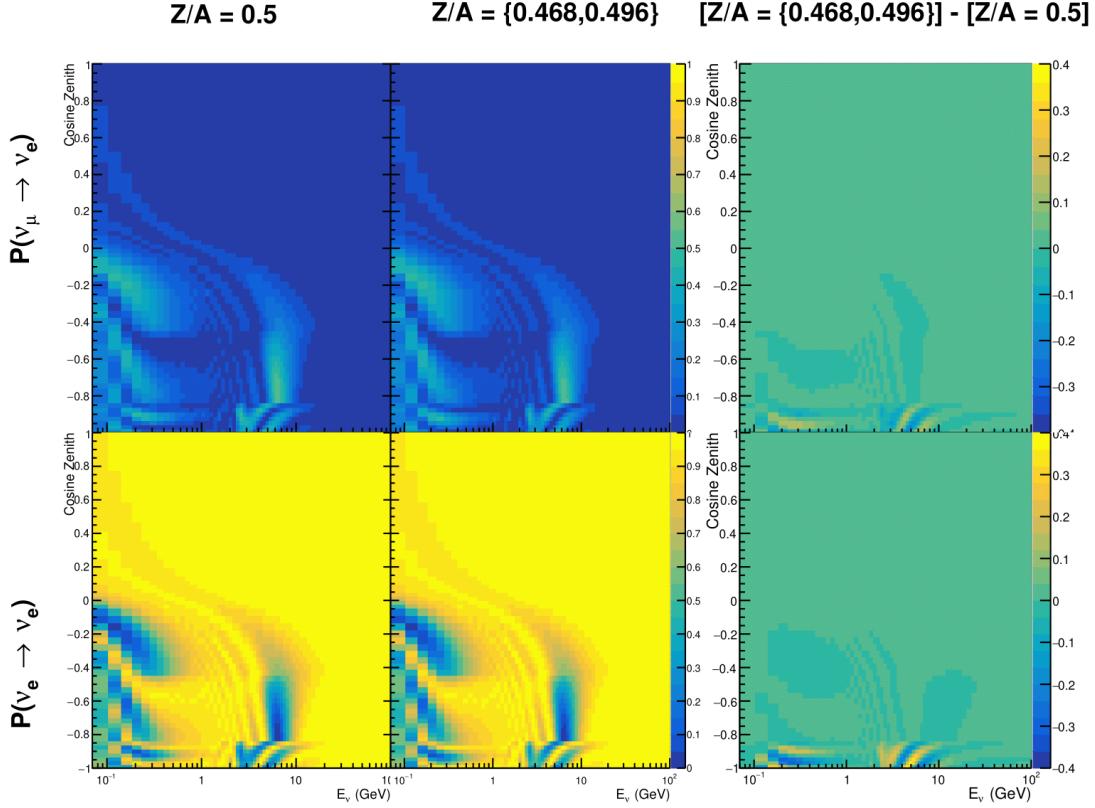


Figure 7.10: The oscillation probability, $P(\nu_\mu \rightarrow \nu_e)$ (top row) and $P(\nu_e \rightarrow \nu_e)$ (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the constant $Z/A = 0.5$ approximation which is used in the official SK-only analysis. The middle column gives the probabilities where $Z/A = [0.468, 0.498]$ values are used, as given in Table 2.3. The right column illustrates the difference in oscillation probability between the two different techniques.

used. The layers of the prem model are combined into four to reduce calculation time, with a quadratic fit to each section. This fit was not performed by the author of the thesis and is documented in [213]. The coefficients of the quadratic fit to each layer are given in Table 7.1 with the final distribution illustrated in Figure 7.11. The quadratic approximation is clearly much closer to the PREM model as compared to the constant density approximation.

The effect of using the quadratic density per $\cos(\theta_Z)$ model is highlighted in Figure 7.12. The slight discontinuity in the oscillation probability around $\cos(\theta_Z) \sim -0.45$ in the fixed density model, which is due to the transition to

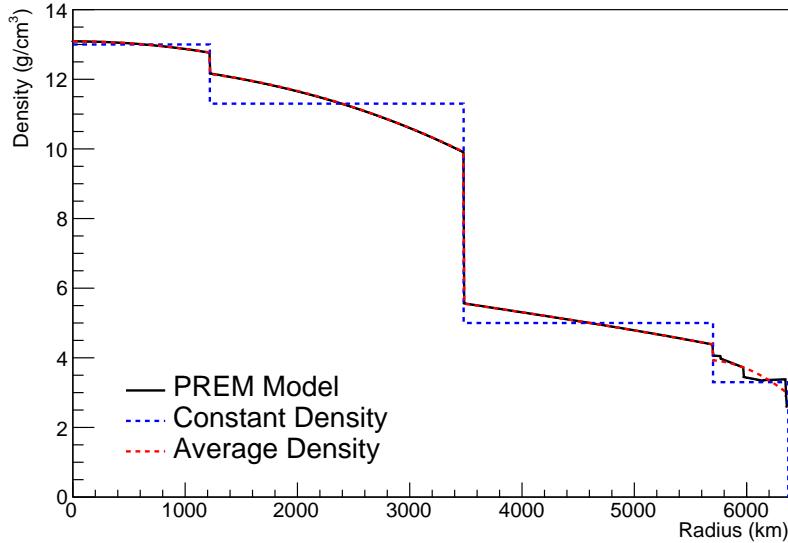


Figure 7.11: The density of the Earth given as a function of the radius, as given by the PREM model (Black), the constant density four-layer approximation (Blue), as used in the official SK-only analysis, and the quadratic approximation of the PREM model (Red).

Layer	Outer Radius [km]	Density [g/cm ³]
Inner Core	1220	$13.09 - 8.84x^2$
Outer Core	3480	$12.31 + 1.09x - 10.02x^2$
Lower Mantle	5701	$6.78 - 1.56x - 1.25x^2$
Transition Zone	6371	$-50.42 + 123.33x - 69.95x^2$

Table 7.1: The quadratic polynomial fits to the PREM model for four assumed layers of the PREM model. The fit to calculate the coefficients is given in [213], where $x = R / R_{\text{Earth}}$.

mantle layer boundary, has been reduced. This is expected as the difference in the density across this boundary is significantly smaller in the quadratic density model as compared to the constant density model. Whilst the difference in density across the other layer transitions is reduced, there is still a significant difference. This means the discontinuities in the oscillation probabilities remain but are significantly reduced. However, as the quadratic density approximation matches the PREM model well in this region, these discontinuities are due to the Earth model rather than an artifact of the oscillation calculation.

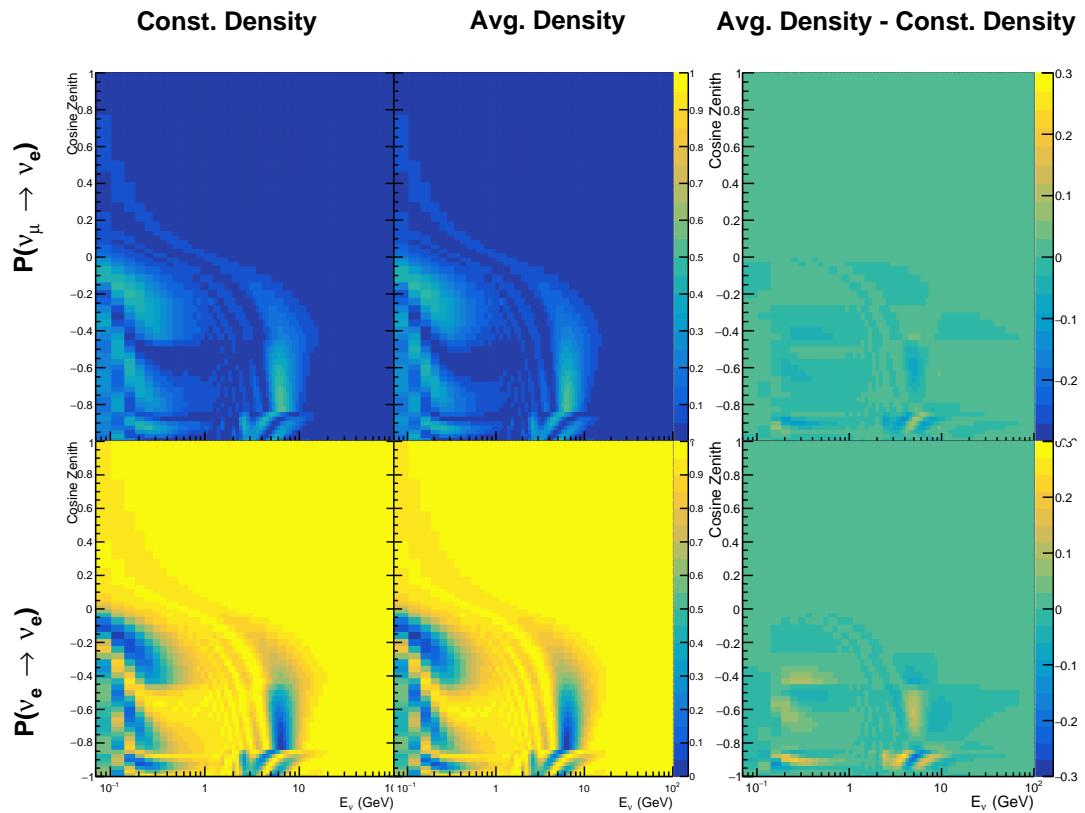


Figure 7.12: The oscillation probability, $P(\nu_\mu \rightarrow \nu_e)$ (top row) and $P(\nu_e \rightarrow \nu_e)$ (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the four-layer constant density approximation is used. The middle column gives the probabilities where the density is integrated over the trajectory, using the quadratic PREM approximation, for each $\cos(\theta_Z)$ is used. The right column illustrates the difference in oscillation probability between the two different techniques.

²⁷⁵⁵ 7.4 Production Height Averaging

²⁷⁵⁶ As discussed in section 2.5, the height at which the cosmic ray flux interacts
²⁷⁵⁷ in the atmosphere is not known on an event-by-event basis. The production
²⁷⁵⁸ height can vary from the Earth’s surface to $\sim 50\text{km}$ above that. The SK-only
²⁷⁵⁹ analysis methodology (described in section 7.1) for including the uncertainty
²⁷⁶⁰ on the production height is to include variations from the Honda model when
²⁷⁶¹ pre-calculating the oscillation probabilities prior to the fit. This technique is not
²⁷⁶² possible for this analysis which uses continuous oscillation parameters that can
²⁷⁶³ not be known prior to the fit. Consequently, an analytical averaging technique
²⁷⁶⁴ was developed in [213]. The author of this thesis was not responsible for the
²⁷⁶⁵ derivation of the technique but has performed the implementation and validation
²⁷⁶⁶ of the technique for this analysis alone.

²⁷⁶⁷ Using the 20 production heights per Monte Carlo neutrino event, provided
²⁷⁶⁸ as 5% percentiles from the Honda flux model, a production height distribution
²⁷⁶⁹ $p_j(h|E_\nu, \cos \theta_Z)$ is built for each neutrino flavour $j = \nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$. In practice, a
²⁷⁷⁰ histogram is filled with 20 evenly spaced bins in production height h between
²⁷⁷¹ 0 and 50km. The neutrino energy and cosine zenith binning of the histogram
²⁷⁷² are the same as that provided in section 7.1. The average production height,
²⁷⁷³ $\bar{h} = \int dh \frac{1}{4} \sum_j p_j(h|E_\nu, \cos(\theta_Z))$, is calculated. This assumes a linear average over
²⁷⁷⁴ the four flavours of neutrino which are considered to be generated in cosmic
²⁷⁷⁵ ray showers. The production height binning of this histogram is then translated
²⁷⁷⁶ into $\delta t(h) = t(\bar{h}) - t(h)$, where $t(x)$ is the distance travelled along the trajectory
²⁷⁷⁷ in the atmosphere from some production height, x .

²⁷⁷⁸ For the i^{th} traversed layer, the transition amplitude, $D_i(t_{i+1}, t_i)$, is computed.
²⁷⁷⁹ The time-ordered product of these is then used as the overall transition amplitude
²⁷⁸⁰ via

$$A(t_{n+1}, t_0) = D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1) D_0(t_1, t_0), \quad (7.6)$$

2781 where,

$$\begin{aligned} D_n(t_{n+1}, t_n) &= \exp[-iH_n(t_{n+1} - t_n)] \\ &= \sum_{k=1}^3 C_k \exp[ia_k(t_{n+1} - t_n)] \end{aligned} \quad (7.7)$$

2782 is expressed as a diagonalised time-dependent solution to the Schrodinger
 2783 equation. The 0th layer is the propagation through the atmosphere and is the
 2784 only term that depends on the production height. Using the substitution $t_0 =$
 2785 $t(\bar{h}) - \delta t(h)$, it can be shown that

$$D_0(t_1, t_0) = D_0(t_1, \bar{h})D_0(\delta t). \quad (7.8)$$

2786 Thus Equation 7.6 becomes

$$\begin{aligned} A(t_{n+1}, t_0) &= D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1)D_0(t_1, \bar{h})D(\delta t) \\ &= A(t_{n+1}, \bar{h}) \sum_{k=1}^3 C_k \exp[ia_k \delta t], \\ &= \sum_{k=1}^3 B_k \exp[ia_k \delta t]. \end{aligned} \quad (7.9)$$

2787 The oscillation probability averaged over production height is then calculated
 2788 as

$$\begin{aligned} \bar{P}(\nu_j \rightarrow \nu_i) &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) P(\nu_j \rightarrow \nu_i) \\ &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) A(t_{n+1}, t_0) A^*(t_{n+1}, t_0) \\ &= \sum_{km} (B_k)_{ij} (B_m)_{ij}^* \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) \exp[i(a_k - a_m) \delta t]. \end{aligned} \quad (7.10)$$

2789 It is important to note that the exact value of \bar{h} used does not matter as the
 2790 values of δt would change to compensate for any modification to the value of \bar{h} .

2791 In practice, implementation in CUDAProb3 [216] is relatively straightforward
 2792 as the majority of these terms are already calculated in the standard oscillation
 2793 calculation. Figure 7.13 illustrates the results of the production height averaging.

2794 As expected, the main effect is observed in the low-energy downward-going
 2795 and horizontal-going events. Upward-going events have to travel the radius
 2796 of the Earth, $R_E = 6371\text{km}$, where the production height uncertainty is a small
 2797 fraction of the total path length.

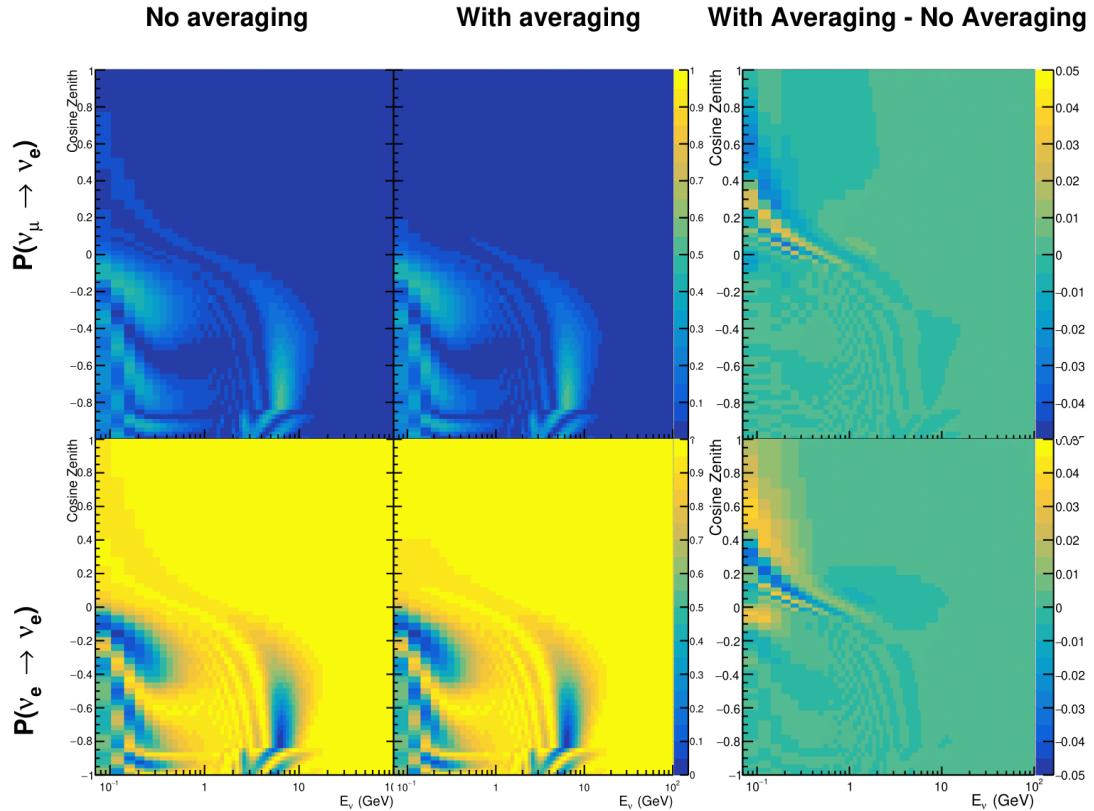


Figure 7.13: The oscillation probability, $P(\nu_\mu \rightarrow \nu_e)$ (top row) and $P(\nu_e \rightarrow \nu_e)$ (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where a fixed production height of 25km is used. The middle column gives the probabilities where the production height is analytically averaged. The right column illustrates the difference in oscillation probability between the two different techniques.

8

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Oscillation Analysis

2800 Using the samples and systematics defined in chapter 6, this chapter documents
2801 a simultaneous beam and atmospheric oscillation analysis from the T2K and SK
2802 experiments. The MaCh3 Bayesian MCMC framework introduced in chapter 4
2803 is used for all studies performed within this thesis.

2804 The MaCh3 framework used throughout this thesis has been validated through
2805 many tests. The code which handles the beam far detector samples was developed
2806 by the author and validated by comparison to the 2020 T2K analysis [75]. The
2807 sample event rates and likelihood evaluations of beam samples generated by the
2808 framework used within this thesis were compared to those from the T2K analysis
2809 by the author of this thesis. Variations of the sample predictions were compared
2810 at $\pm 1\sigma$ and $\pm 3\sigma$ and good agreement was found in all cases. A similar study, led
2811 by Dr. C. Wret was used to validate the near detector portion of the code [206].
2812 The implementation of the atmospheric samples within MaCh3 was completed
2813 and cross-checked by the author of this thesis against the P-Theta framework
2814 (Introduced in section 3.2). Both fitters are provided with the same inputs and can
2815 therefore cross-validate each other. These validations compared the event rate and
2816 likelihood calculation. Documentation of all the above validations can be found
2817 in [206]. These stringent validations ensure that the code is doing as intended.

2818 8.1 Monte Carlo Prediction

2819 Using the three sets of dial values (generated, pre-fit, and post-fit tunes) defined in
 2820 subsection 6.4.3, the predicted event rates for each sample are given in Table 8.1.
 2821 The oscillated (AsimovA defined in Table 2.2) and un-oscillated event rates
 2822 are calculated for each tune.

Sample	Total Predicted Events					
	Generated		Pre-fit		Post-fit	
	Osc	UnOsc	Osc	UnOsc	Osc	UnOsc
SubGeV- <i>elike</i> -0dcy	7121.0	7102.6	6556.8	6540.0	7035.2	7015.7
SubGeV- <i>elike</i> -1dcy	704.8	725.5	693.8	712.8	565.7	586.0
SubGeV- <i>mulike</i> -0dcy	1176.5	1737.2	1078.6	1588.1	1182.7	1757.1
SubGeV- <i>mulike</i> -1dcy	5850.7	8978.1	5351.7	8205.1	5867.0	9009.9
SubGeV- <i>mulike</i> -2dcy	446.9	655.2	441.6	647.7	345.9	505.6
SubGeV- <i>pi0like</i>	1438.8	1445.4	1454.9	1461.1	1131.1	1136.2
MultiGeV- <i>elike</i> -nue	201.4	195.6	201.1	195.3	202.6	196.7
MultiGeV- <i>elike</i> -nuebar	1141.5	1118.3	1060.7	1039.5	1118.5	1095.7
MultiGeV- <i>mulike</i>	1036.7	1435.8	963.1	1334.1	1015.2	1405.9
MultiRing- <i>elike</i> -nue	1025.1	982.2	1026.8	984.3	1029.8	986.4
MultiRing- <i>elike</i> -nuebar	1014.8	984.5	991.0	962.0	1008.9	978.5
MultiRing- <i>mulike</i>	2510.0	3474.4	2475.6	3425.8	2514.6	3480.4
MultiRingOther-1	1204.5	1279.1	1205.8	1280.3	1207.4	1281.0
PCStop	349.2	459.2	338.4	444.7	346.8	456.1
PCThru	1692.8	2192.5	1661.5	2149.8	1689.2	2187.8
UpStop-mu	751.2	1295.0	739.7	1271.6	750.4	1293.0
UpThruNonShower-mu	2584.4	3031.6	2577.9	3019.4	2586.8	3034.0
UpThruShower-mu	473.0	488.6	473.2	488.7	473.8	489.4
FHC1Rmu	328.0	1409.2	301.1	1274.7	345.1	1568.0
RHC1Rmu	133.0	432.3	122.7	396.2	135.0	443.9
FHC1Re	84.6	19.2	77.4	18.2	93.7	19.7
RHC1Re	15.7	6.4	14.6	6.1	15.9	6.3
FHC1Re1de	10.5	3.2	10.3	3.1	8.8	2.9

Table 8.1: The Monte Carlo predicted event rate of each far detector sample used within this analysis. Three model parameter tunes are considered, as defined in subsection 6.4.3. Un-oscillated and oscillated predictions are given, where the oscillated predictions assume Asimov A oscillation parameters provided in Table 2.2.

2823 Generally, the samples which target CCQE interaction modes observe a
 2824 decrease in prediction when using the pre-fit dial values. This is in accordance
 2825 with the Monte Carlo being produced assumed $M_A^{QE} = 1.21\text{GeV}$ [164] whilst the

2826 pre-fit dial value is set to $M_A^{QE} = 1.03\text{GeV}$ as suggested by [204]. Furthermore,
 2827 the predicted event rates of samples that target CCRES interaction modes are
 2828 significantly reduced when considering the post-BANFF fit. This follows the ob-
 2829 servations in subsection 6.4.3. The strength of the accelerator neutrino experiment
 2830 can be seen in the remarkable difference between the oscillated and unoscillated
 2831 predictions in the FHC1Rmu and RHC1Rmu samples. There is a very clear decrease
 2832 in the expected event rate between the oscillated and un-oscillated predictions
 2833 which is not as obvious as in the atmospheric samples. This is due to the fact that
 2834 the beam energy is tuned to the maximum disappearance probability, which is
 2835 not the case for the naturally generated atmospheric neutrinos.

2836 8.1.1 Likelihood Scans

2837 Using the definition of the likelihood presented in section 6.5, the contribution
 2838 of each sample to the likelihood from a variation of a particular parameter can
 2839 be studied. This process identifies which samples drive the determination of the
 2840 oscillation parameters in the joint fit. Figure 8.1 presents the variation of all the
 2841 samples (beam and atmospheric) at the far detector to the oscillation parameters
 2842 of interest: δ_{CP} , $\sin^2(\theta_{13})$, $\sin^2(\theta_{23})$, and Δm_{32}^2 . These plots are colloquially called
 2843 ‘likelihood scan’ (or ‘log-likelihood scans’). The process of making these plots is
 2844 as follows. An Asimov data set (following the technique detailed in section 6.5) is
 2845 built using the AsimovA oscillation parameters and pre-fit systematic tune. The
 2846 Monte Carlo is then reweighted using the value of the oscillation parameter at
 2847 each point on the x-axis of the scan. The likelihood is then calculated between
 2848 the Asimov data and Monte Carlo prediction and plotted.

2849 Due to the caveat of fixed systematic parameters and the correlations between
 2850 oscillation parameters being ignored when creating these likelihood scans, the
 2851 value of $\chi^2 = 1$ (or $-2 \times \ln(\text{Likelihood}) = 1$) does not equate to the typical
 2852 1σ sensitivity. However, it does give an indication of which samples respond
 2853 mostly strongly to variations in a particular oscillation parameter. The point at

2854 which the likelihood tends to zero illustrates the value of the parameter used
 2855 to build the Asimov data prediction.

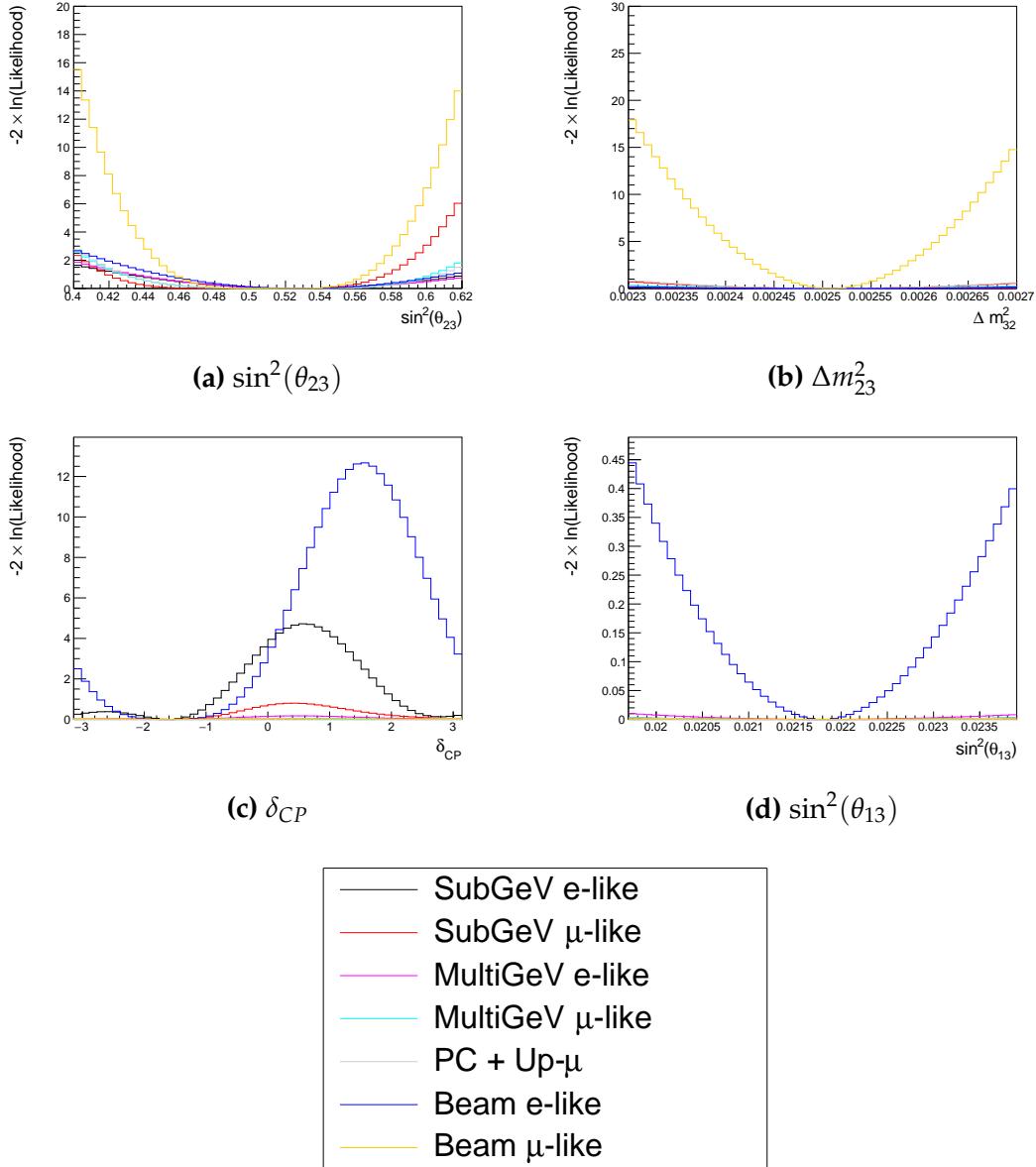


Figure 8.1: The response of the likelihood, as defined in section 6.5, illustrating the response of the samples to a variation of an oscillation parameter.

2856 The sensitivity to $\sin^2(\theta_{23})$ is mostly dominated by the beam muon-like
 2857 samples. The response of an individual atmospheric sample is small but non-
 2858 negligible such that the summed response over all atmospheric samples becomes
 2859 comparable to that of the muon-like beam samples. Consequently, the sensitivity

of the joint fit to $\sin^2(\theta_{23})$ would be expected to be greater than the beam-only analysis. The only sample which responds to the $\sin^2(\theta_{13})$ oscillation parameter is the electron-like beam samples. Consequently, no increase in sensitivity beyond that of the T2K-only analysis would be expected from the joint fit. Regardless, the sensitivity of the beam sample is significantly weaker than the external reactor constraint so prior knowledge will dominate any sensitivity to $\sin^2(\theta_{13})$ which is included within this thesis. The Δm_{21}^2 and $\sin^2(\theta_{12})$ parameters are not considered as there is simply no sensitivity in any sample considered within this analysis. The response to Δm_{32}^2 is completely dominated by the beam muon-like samples. This is because the beam neutrino energy is specifically tuned to match the maximal disappearance probability. Despite this, improvements to the $|\Delta m_{32}^2|$ sensitivity may be expected due to additional mass hierarchy determination added by the atmospheric samples.

Two-dimensional scans of the appearance ($\sin^2(\theta_{13})$ - δ_{CP}) and disappearance ($\sin^2(\theta_{23})$ - Δm_{32}^2) parameters are illustrated in Figure 8.2 and Figure 8.3, respectively. The caveat of fixed systematic parameters and correlations between other oscillation parameters being neglected still apply.

The appearance log-likelihood scans show the distinct difference in how the beam and atmospheric samples respond. The beam samples have an approximately constant width of the 2σ and 3σ contours, throughout all ranges of δ_{CP} . Whereas, the response of the atmospheric samples to $\sin^2(\theta_{13})$ is very strongly correlated to the value of δ_{CP} . At higher values of $\sin^2(\theta_{13})$, two lobes appear around $\delta_{CP} \sim -\pi/2$ and $\delta_{CP} \sim 2.4$. Consequently, this difference allows some of the degeneracy in a beam-only fit to be broken. Comparing the beam-only and joint fit likelihood scans, the 2σ continuous contour in δ_{CP} for beam samples becomes closed when the atmospheric samples are added. This may result in a stronger sensitivity to δ_{CP} . Similarly, the width of the 3σ contours also becomes dependent upon the value of δ_{CP} . Furthermore, atmospheric samples have little sensitivity to $\sin^2(\theta_{13})$ on their own, as evidenced in Figure 8.1, but may improve sensitivity to the parameter when combined within the simultaneous

2890 fit. It is important to remember that these likelihood scans are not sensitivity
2891 measurements as the systematic parameters are fixed and the correlation between
2892 oscillation parameters is neglected. However, they are a very encouraging result
2893 for the joint fit.

2894 The disappearance log-likelihood scans in $\sin^2(\theta_{23})$ - Δm_{32}^2 space (Figure 8.3)
2895 show the expected behaviour when considering the one-dimensional scans
2896 already discussed. The uncertainty on the width of $|\Delta m_{32}^2|$ is mostly driven by the
2897 beam samples. However, the width of this contour in the inverted mass region
2898 ($\Delta m_{32}^2 < 0$) is significantly reduced due to the ability of the atmospheric samples
2899 to select the correct (normal) mass hierarchy. The width of the uncertainty
2900 in $\sin^2(\theta_{23})$ is also reduced compared to the beam-only sensitivities, with a
2901 further decrease in the inverted hierarchy region due to the better mass hierarchy
2902 determination.

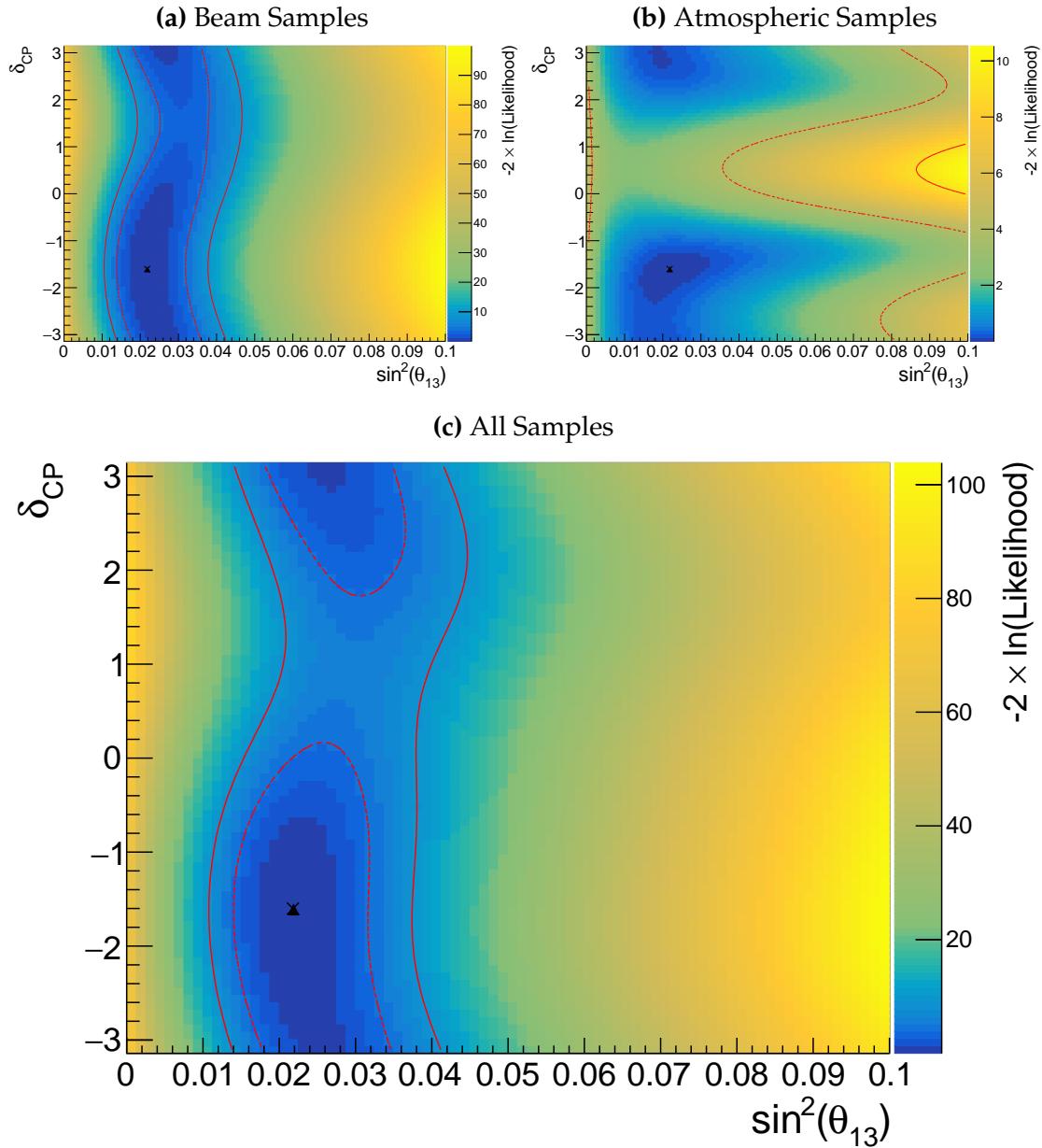


Figure 8.2: Two-dimensional log-likelihood scan of the appearance ($\sin^2(\theta_{13})$ - δ_{CP}) parameters showing the response of the beam samples (top left), atmospheric samples (top right) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are known to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.

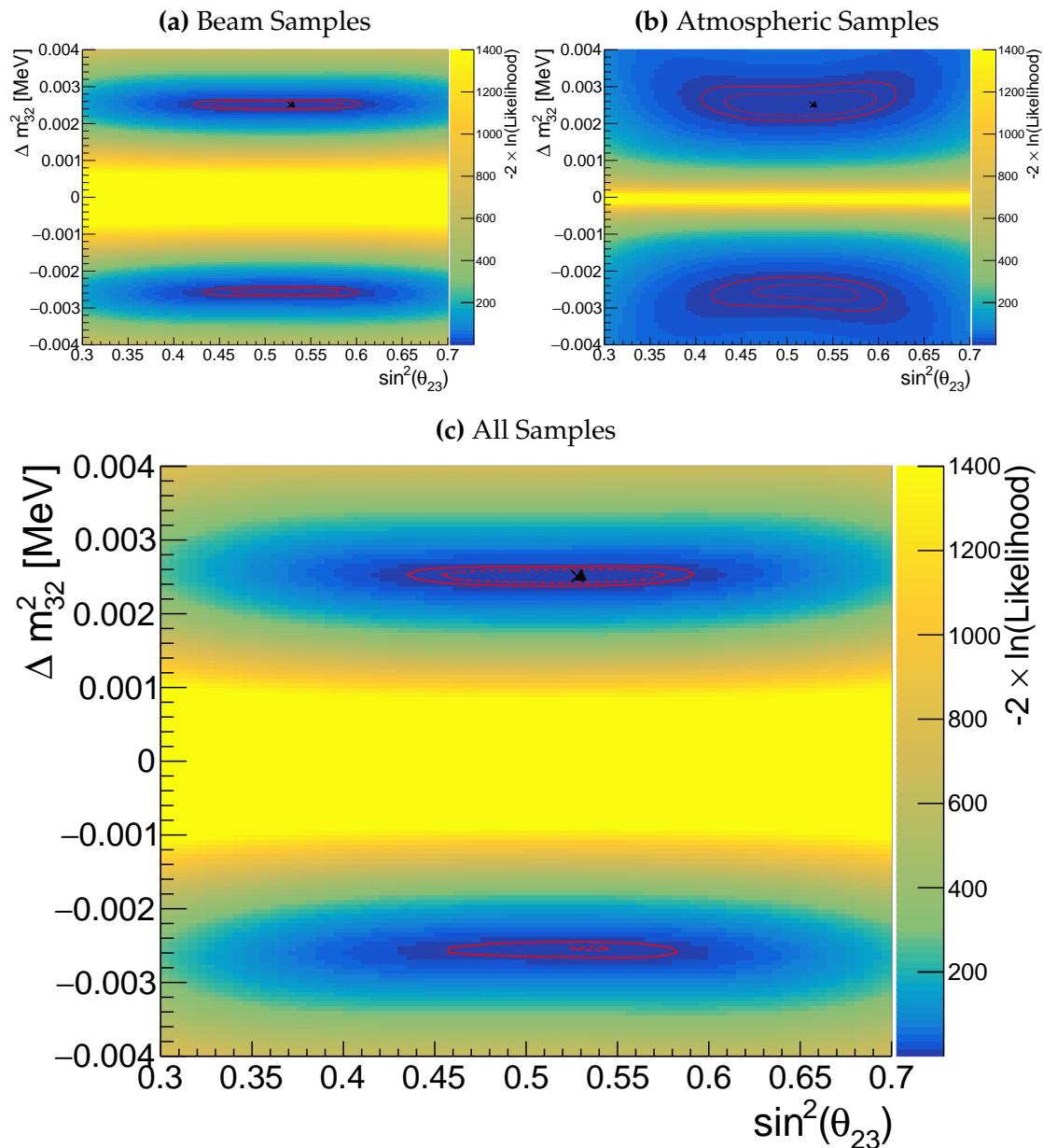


Figure 8.3: Two-dimensional log-likelihood scan of the disappearance ($\sin^2(\theta_{23})$)- Δm_{32}^2) parameters showing the response of the beam samples (top left), atmospheric samples (top right) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are known to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.

The likelihood scans illustrated thus far only consider the sensitivity of this analysis for a fixed set of true oscillation parameters, namely Asimov A defined in Table 2.2. Whilst computationally infeasible to run many fits at different parameter sets, it is possible to calculate the likelihood response to different Asimov data sets. Figure 8.4 and Figure 8.5 illustrate how the sensitivity changes for differing true values of δ_{CP} and $\sin^2(\theta_{23})$, respectively. For both of these plots, the other oscillation parameters are fixed at their Asimov A values. Consequently, the caveat of fixed systematic parameters and correlations between other oscillation parameters being neglected still applies.

To explain how these plots are made, consider Figure 8.4. This plot is built by considering multiple one-dimensional log-likelihood scans, each creating an Asimov data set with the value of δ_{CP} taken from the x-axis. The likelihood to this particular Asimov data set is calculated after reweighting the Monte Carlo prediction to each value of δ_{CP} on the y-axis.

Figure 8.4 illustrates the sensitivity to δ_{CP} . To interpret this plot, larger contours results in more phase space being excluded from the 1σ region. The 1σ intervals contain regions where the beam and atmospheric samples have discontinuous contours. For example, for the x-axis value of $\delta_{CP} = 0$, the beam samples sensitivity would include two discontinuous regions excluded from the 1σ interval: $\delta_{CP} \sim 0$ and $\delta_{CP} \sim \pi$. This behaviour is also seen in atmospheric samples response but at a value of $\delta_{CP} \sim -1$. This difference allows the joint fit to have increased sensitivity to these regions. Consequently, the difference between the beam-only and joint beam-atmospheric fit should be studied using multiple Asimov data sets.

Despite the increased sensitivity at 1σ , the 2σ intervals from the joint fit are more similar to the two independent sensitivities and the off-diagonal degeneracies mostly remain. This indicates that the joint fit has the strength to aid parameter determination but can not entirely break the degeneracies in δ_{CP} at higher confidence levels.

Figure 8.5 illustrates a similar analysis as above, although the value of $\sin^2(\theta_{23})$ is varied and δ_{CP} is fixed to the Asimov A parameter value. Due to the beam parameters and baseline being tuned to specifically target this oscillation parameter, the average sensitivity of the beam samples is stronger than the atmospheric samples. However, the degeneracy around maximal mixing ($\sin^2(\theta_{23}) = 0.5$) is significantly more peaked in the beam samples compared to the atmospheric samples. This means that a value of $\sin^2(\theta_{23}) \sim 0.56$ would be contained within the 1σ confidence interval for a true value of $\sin^2(\theta_{23}) \sim 0.46$ if using the beam-only analysis, whereas it would be excluded in the joint analysis.

This behaviour is strengthened when considering the 2σ intervals, to the point where two distinct discontinuous regions of the 2σ intervals exist around the Asimov point $\sin^2(\theta_{23}) \sim 0.41, 0.6$. Given the caveat of only considering likelihood scans, the joint analysis would mostly eliminate the discontinuous intervals in these regions. This means that the joint fit could feasibly have an increased preference for the correct octant hypothesis.

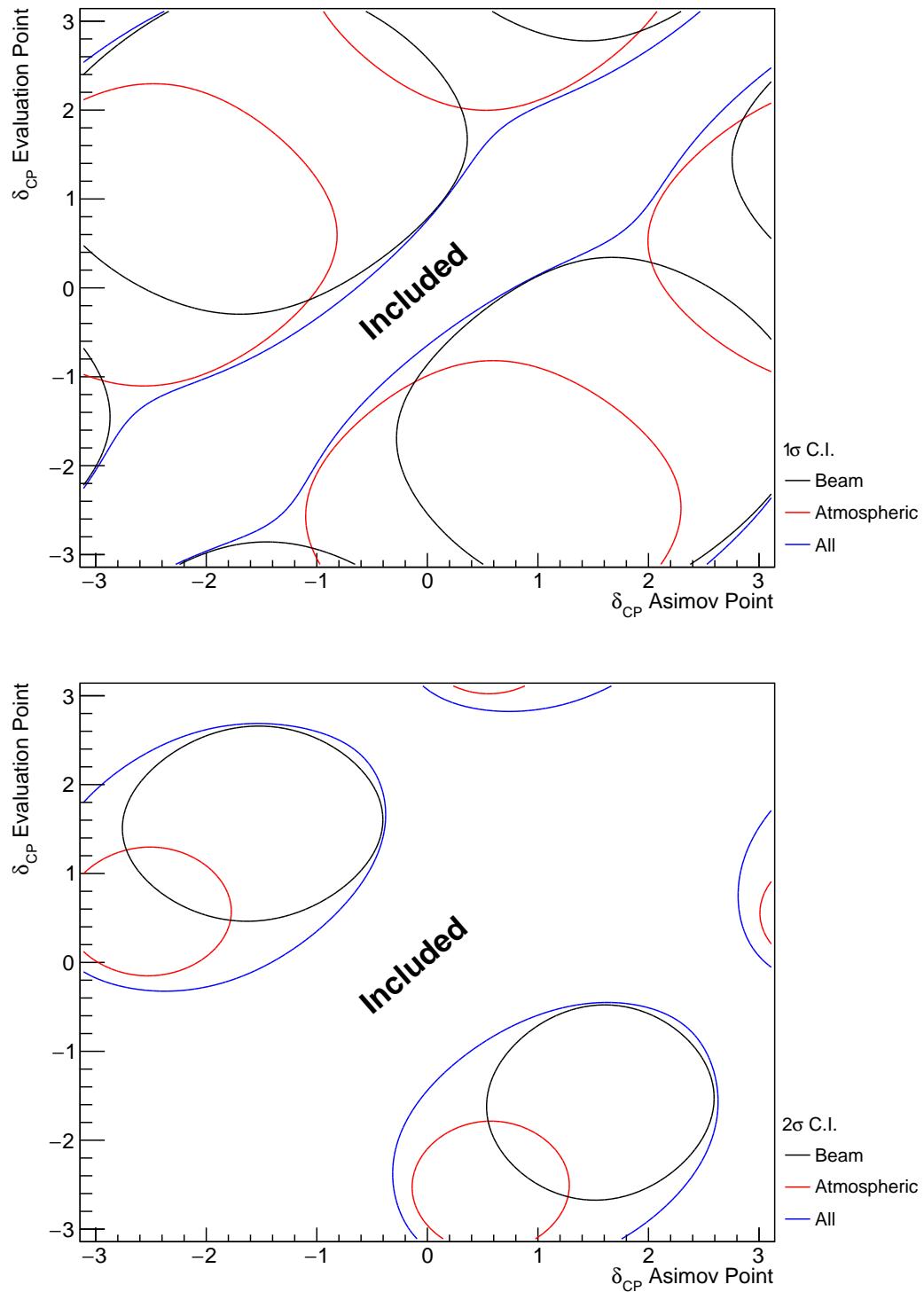


Figure 8.4: A series of one-dimensional likelihood scans over δ_{CP} , where an Asimov data set is built for each value of δ_{CP} on the x-axis and the likelihood is evaluated for each value of δ_{CP} on the y-axis. The diagonal represents the minimum log-likelihood and defines the region included within the 1σ (Top) and 2σ (Bottom) confidence intervals. The beam (black) and atmospheric (red) samples are individually plotted and the joint fit (blue) is the sum of the two.

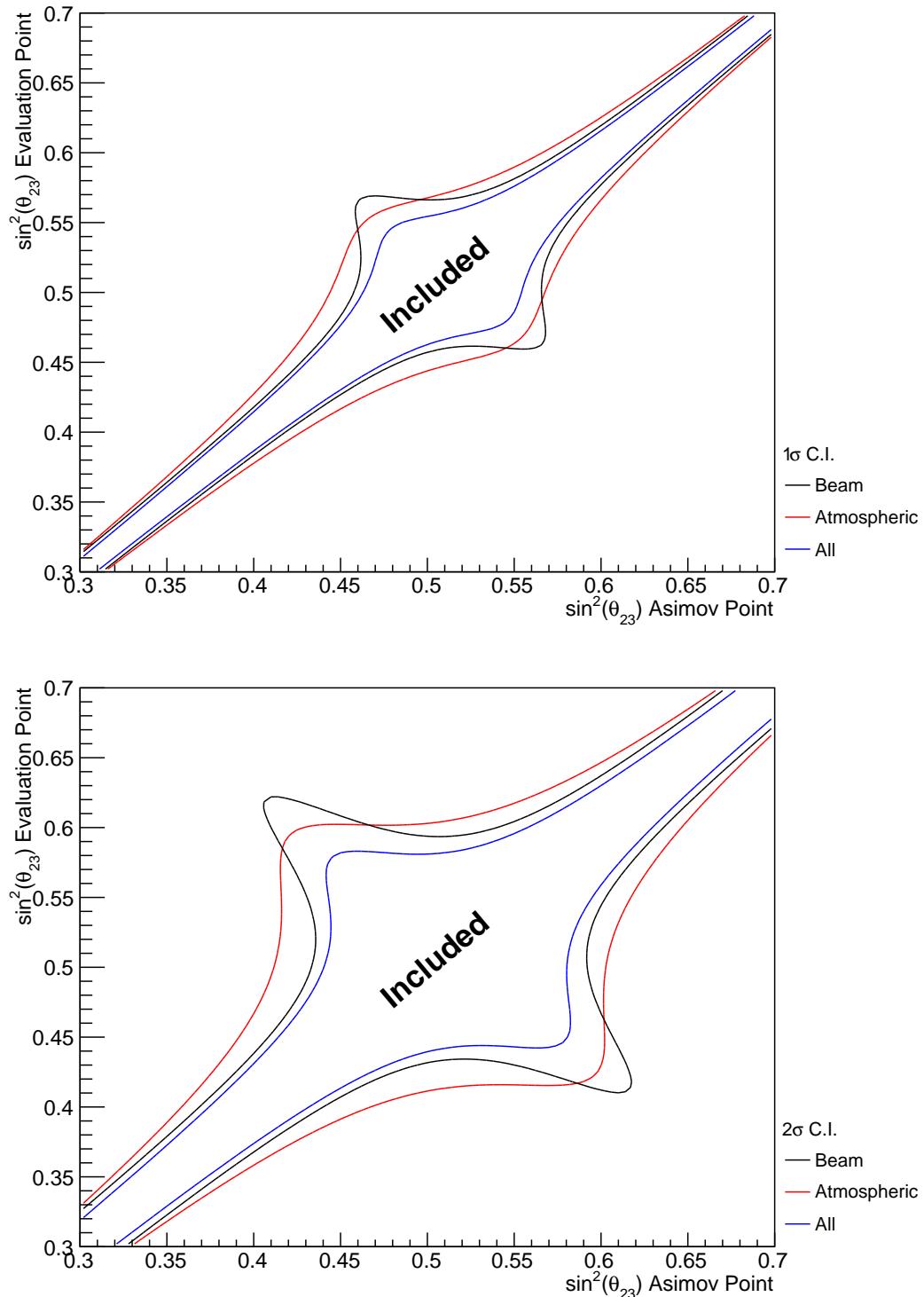


Figure 8.5: A series of one-dimensional likelihood scans over $\sin^2(\theta_{23})$, where an Asimov data set is built for each value of $\sin^2(\theta_{23})$ on the x-axis and the likelihood is evaluated for each value of $\sin^2(\theta_{23})$ on the y-axis. The diagonal represents the minimum log-likelihood and defines the region included within the 1σ (Top) and 2σ (Bottom) confidence intervals. The beam (black) and atmospheric (red) samples are individually plotted and the joint fit (blue) is the sum of the two.

Alongside oscillation parameters (Figure 8.1), the sensitivity to systematic parameters can also be studied for the joint fit. As some of these parameters are correlated between the beam and atmospheric events, the response of the atmospheric samples can modify the constraint. This means the systematics can have additional constraints than they would from a beam-only analysis. Therefore, the response from the beam and the atmospheric samples to various systematic parameters has been compared in Figure 8.6. The Asimov data set has been created using the AsimovA oscillation parameter and the pre-fit systematic tune. For example, the systematic parameter controlling the effective axial mass coupling in CCQE interactions, M_A^{QE} , is clearly dominated by the ND constraint. An example where the response of the atmospheric sample is approximately similar to the near detector constraint is the 2p2h CtoO normalisation systematic. This systematic models the scaling of the 2p2h interaction cross-section on a carbon target to an oxygen target. There are also systematics which have no near detector constraint. For example, the systematic parameters which describe the normalisation of the NC1Gamma and NCOther interaction modes. The atmospheric and beam samples can have similar sensitivity to these systematics due to their similar composition in energy and interaction mode. An example of how the atmospheric samples can help constrain systematic parameters used within the T2K-only analysis, these NC background events in beam electron-like samples will be more constrained with the additional sensitivity of atmospheric samples. This would be expected to reduce the overall uncertainty of the beam electron-like event rates in the joint analysis compared to the beam-only studies. This could modify the sensitivity of the beam samples due to the more constrained background events.

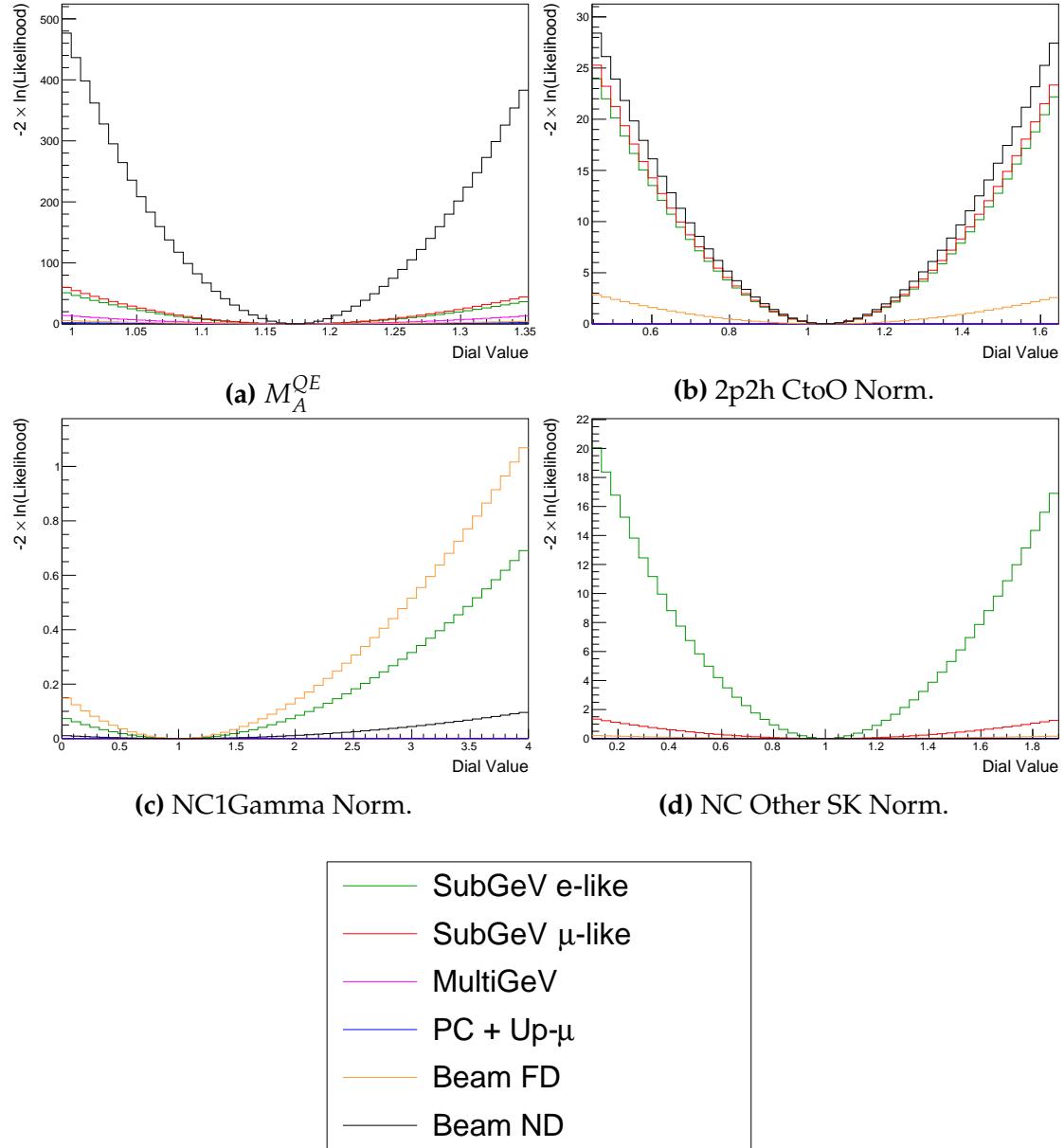


Figure 8.6: The response of the likelihood, as defined in section 8.2, illustrating the response of the samples to the various cross-section systematic parameters.

2972 8.2 Sensitivities

2973 DB: Give HPD values as a bin width

2974 DB: Remove OA2020 from plots

2975 The sensitivities of the joint T2K and SK oscillation analysis are presented
2976 in the form of Asimov fits. These fits consider beam samples from the near
2977 and far detector alongside atmospheric samples at SK. This technique builds an
2978 Asimov data set (following section 6.5) using the AsimovA oscillation parameters
2979 and post-BANFF systematic tune, which is then ffit. This technique eliminates
2980 statistical fluctuations from the data therefore providing the maximum sensitivity
2981 of the analysis.

2982 In practice, the Asimov fits presented within this analysis are modified from
2983 the above definition. An Asimov prediction of both beam and atmospheric far
2984 detector samples is fit whilst the true data is used for near detector samples. The
2985 Asimov predictions at the far detector are built using the post-BANFF tune (as
2986 discussed in section 3.2). These modifications mean that the results are equivalent
2987 to performing a far detector Asimov fit using inputs from the BANFF data fit.
2988 Consequently, this allows the results to be cross-checked with the results from
2989 the P-Theta analysis. The comparison has been performed and is documented in
2990 [222]. No significant discrepancies were found between the fitters.

2991 This section proceeds with the following studies. Firstly, the sensitivity
2992 of the atmospheric samples using the correlated detector model is detailed in
2993 subsection 8.2.1. This includes studying the choice of applying the 2020 PDG
2994 reactor constraint [76] to the atmospheric samples, which is documented in
2995 subsection 8.2.2. Additionally, the effect of applying the near-detector constraints
2996 onto the atmospheric samples is discussed in subsection 8.2.3. The main result is
2997 the sensitivity of the simultaneous beam and atmospheric fit. The sensitivities,
2998 both with and without the application of the reactor constraint, are presented
2999 in subsection 8.2.4 and subsection 8.2.5, respectively. To indicate the benefit
3000 of the joint analysis, the sensitivities are compared to the 2020 T2K beam-only

sensitivities [75, 223] in subsection 8.2.6 and subsection 8.2.7. The T2K analysis is used as a reference as it uses the same samples and a similar systematic model. As shown in subsection 8.1.1, the response of the beam and atmospheric samples change depending upon the true set of oscillation parameters assumed. Therefore, subsection 8.2.8 documents the sensitivities at an alternative oscillation parameter set. It is important to note that these results have been presented at the Neutrino 2022 conference on behalf of the T2K and SK collaborations [111].

8.2.1 Atmospheric-Only Sensitivity Without Reactor Constraint

This section presents the results of an Asimov fit using samples from the near detector and only atmospheric samples from the far detector. The results are presented as one-dimensional or two-dimensional histograms which have been marginalised over all other parameters using the technique outlined in subsection 4.3.1. Each histogram displays the posterior probability density and illustrates the credible intervals, calculated using the technique in subsection 4.3.2. For this fits, a flat prior is used for $\sin^2(\theta_{13})$ such the reactor constraint is not applied. The Asimov data is generated assuming the AsimovA oscillation parameter set defined in Table 2.2 and the post-BANFF systematic parameter tune.

Figure 8.7 illustrates the posterior probability density for δ_{CP} , marginalised over both hierarchies. The fit favours the known oscillation parameter ($\delta_{CP} = -1.601$) although the posterior probability is very flat through the range of $-\pi < \delta_{CP} < -1$ and $2 < \delta_{CP} < \pi$. There is also a region around $\delta_{CP} \sim 0.4$ which is disfavoured at 2σ . This indicates that the SK samples can rule out some parts of the CP conserving parameter space reasonably well, near $\delta_{CP} \sim 0.4$, when the true value of $\delta_{CP} \sim \pi/2$.

The posterior probability density in Δm_{32}^2 is given in Figure 8.8. This distribution includes steps in both the normal hierarchy (NH, $\Delta m_{32}^2 > 0$) and the inverse hierarchy (IH, $\Delta m_{32}^2 < 0$). The highest posterior probability density is found within the NH, which agrees with the known oscillation parameter value. However, all of the credible intervals span both of the hierarchy hypotheses.

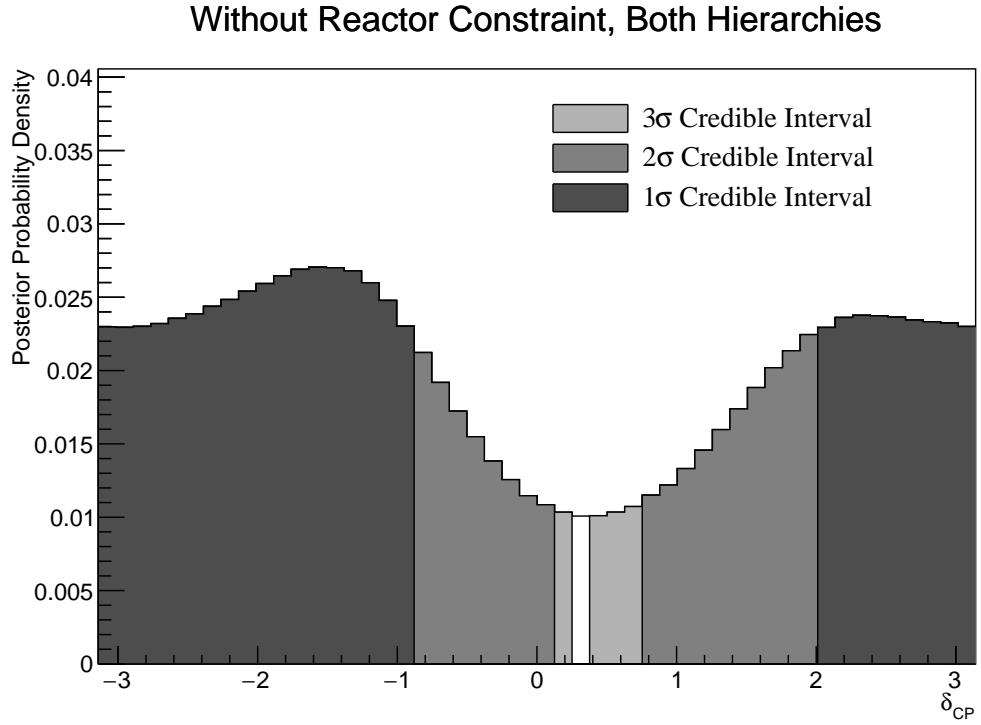


Figure 8.7: The one-dimensional posterior probability density distribution in δ_{CP} , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3030 The known oscillation parameter is $2.509 \times 10^{-3}\text{eV}^2$, which is contained within
 3031 the 1σ credible interval.

	LO ($\sin^2 \theta_{23} < 0.5$)	UO ($\sin^2 \theta_{23} > 0.5$)	Sum
NH ($\Delta m_{32}^2 > 0$)	0.17	0.40	0.58
IH ($\Delta m_{32}^2 < 0$)	0.13	0.29	0.42
Sum	0.31	0.69	1.00

Table 8.2: The distribution of steps in an SK atmospheric-only fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as $B(\text{NH}/\text{IH}) = 1.37$ and $B(\text{UO}/\text{LO}) = 2.24$.

3032 Following the discussion in subsection 4.3.3, the Bayes factor for hierarchy
 3033 preference can be calculated by determining the fraction of steps that fall into the
 3034 NH and the IH regions, as an equal prior is placed on both hypotheses. A similar
 3035 calculation can be performed by calculating the fraction of steps which fall in

Without Reactor Constraint, Both Hierarchies

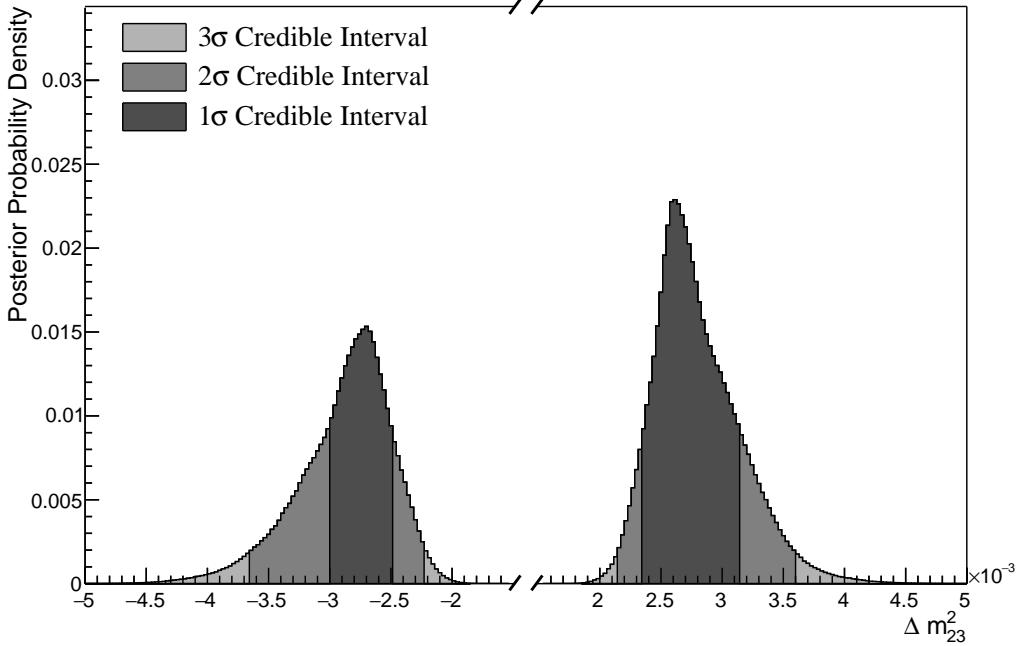


Figure 8.8: The one-dimensional posterior probability density distribution in Δm_{32}^2 , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

the lower octant (LO, $\sin^2 \theta_{23} < 0.5$) or upper octant (UO, $\sin^2 \theta_{23} > 0.5$). The fraction of steps, broken down by hierarchy and octant, are given in Table 8.2. The Bayes factor for preferred hierarchy model is $B(\text{NH}/\text{IH}) = 1.37$. Jeffrey's scale, given in Table 4.1, states this value of the Bayes factor indicates a weak preference for the normal hierarchy model. The Bayes factor for choice of octant is $B(\text{UO}/\text{LO}) = 2.24$. This is also classified as a weak preference for the UO. Both of these show that the fit is returning the correct choice of models (NH and UO) for the known Asimov A oscillation parameters defined in Table 2.2.

The 1 σ credible intervals, broken down by hierarchy, and position in parameter space of the highest posterior probability density is given in Table 8.3. These are taken from the one-dimensional projections of the oscillation parameters, marginalised over all other parameters within the fit. For the known Asimov value of $\delta_{CP} = -1.601$, the 1 σ credible interval rules out a region between $\delta_{CP} = -0.86$ and $\delta_{CP} = 1.96$, when marginalising over both hierarchies. The position

Parameter	Interval	HPD
δ_{CP} , (BH)	$[-\pi, -0.86], [1.96, \pi]$	-1.57
δ_{CP} , (NH)	$[-\pi, -0.86], [1.88, \pi]$	-1.57
δ_{CP} , (IH)	$[-\pi, -0.94], [1.96, \pi]$	-1.57
Δm_{32}^2 (BH) [$\times 10^{-3}\text{eV}^2$]	$[-3.00, -2.50], [2.35, 3.15]$	2.65
Δm_{32}^2 (NH) [$\times 10^{-3}\text{eV}^2$]	$[2.39, 3.04]$	2.64
Δm_{32}^2 (IH) [$\times 10^{-3}\text{eV}^2$]	$[-3.15, -2.45]$	-2.70
$\sin^2(\theta_{23})$ (BH)	$[0.476, 0.59]$	0.542
$\sin^2(\theta_{23})$ (NH)	$[0.476, 0.59]$	0.554
$\sin^2(\theta_{23})$ (IH)	$[0.476, 0.59]$	0.542

Table 8.3: The position of the highest posterior probability density (HPD) and width of the 1σ credible interval for the SK atmospheric-only fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

of the highest posterior density is $\delta_{CP} = -1.57$ which is clearly compatible with the known oscillation parameter value.

The sensitivity of the atmospheric samples to $\sin^2(\theta_{13})$ is presented in Figure 8.9. The likelihood scans presented in Figure 8.1 suggest that the sensitivity to $\sin^2(\theta_{13})$ will be small. This behaviour is also seen in the fit results, where the width of the 1σ credible intervals span the region of $\sin^2(\theta_{13}) = [0.008, 0.08]$. This is more than an order of magnitude worse than the constraint from reactor experiments [76].

As previously discussed, the correlations between oscillation parameters are also important to understand how the atmospheric samples respond. Figure 8.10 illustrates the two dimensional $\sin^2(\theta_{13}) - \delta_{CP}$ sensitivity, marginalised over all other parameters. The shape of the 1σ credible interval shows that the constraining power of the fit on δ_{CP} is dependent upon the value of $\sin^2(\theta_{13})$. Furthermore, they show a strong resemblance to the likelihood scans illustrated in Figure 8.2. Whilst the atmospheric samples do not strongly constrain the

Without Reactor Constraint, Both Hierarchies



Figure 8.9: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{13})$, marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3065 value of $\sin^2(\theta_{13})$, the value of $\sin^2(\theta_{13})$ does impact the atmospheric samples'
 3066 sensitivity to δ_{CP} .

3067 The $\sin^2(\theta_{23}) - \Delta m_{32}^2$ disappearance contours are illustrated in Figure 8.11.
 3068 As expected, the area contained in the inverted hierarchy 1σ credible interval is
 3069 slightly smaller than that in the normal hierarchy. This follows from the Bayes
 3070 factor showing a weak preference for NH meaning that more of the steps will exist
 3071 in the $\Delta m_{32}^2 > 0$ region. The known oscillation parameters of $\sin^2(\theta_{23}) = 0.528$
 3072 and $\Delta m_{32}^2 = 2.509 \times 10^{-3} \text{ eV}^2$ are contained within the 1σ credible interval.

3073 Figure 8.12 illustrates the two-dimensional projections for each permutation of
 3074 oscillation parameters which this analysis is sensitive to: δ_{CP} , $\sin^2(\theta_{13})$, $\sin^2(\theta_{23})$,
 3075 and Δm_{32}^2 . The purpose of this plot is to illustrate the correlations between
 3076 the oscillation parameters. The contours are calculated whilst marginalising
 3077 over both hierarchies, however, only the NH is illustrated when plotting the
 3078 Δm_{32}^2 parameter. As expected the correlations play a significant role in these

Without Reactor Constraint, Both Hierarchies

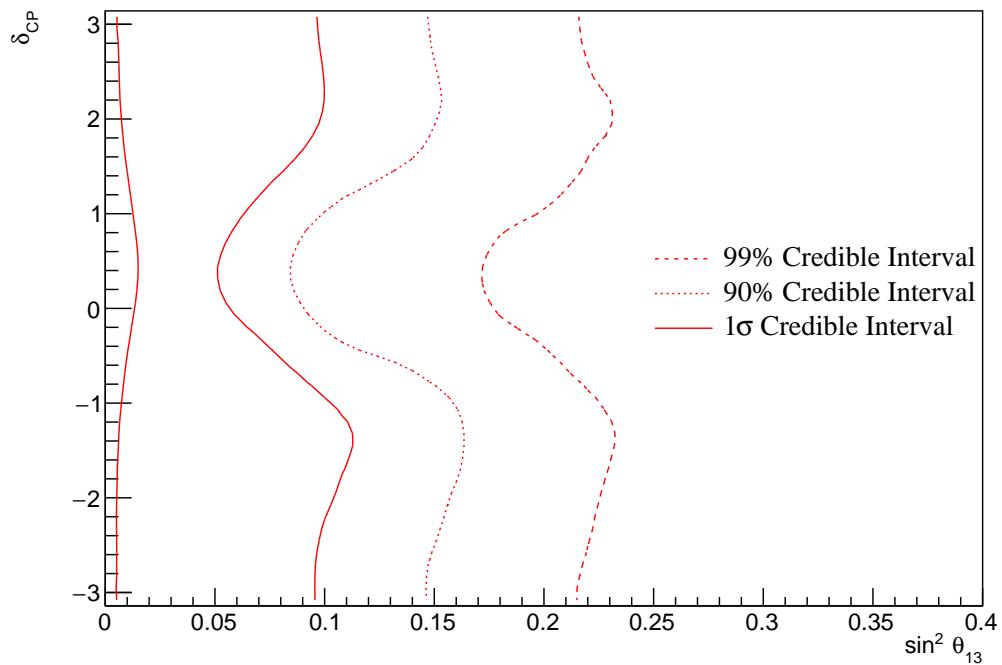


Figure 8.10: The two-dimensional posterior probability density distribution in δ_{CP} – $\sin^2(\theta_{13})$, marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3079 sensitivity measurements, especially the choice of the $\sin^2(\theta_{13})$ constraint. Most
 3080 notably, the application of reactor constraint would be expected to alter both the
 3081 width and position of the Δm_{32}^2 intervals due to the strong correlation between
 3082 the parameters.

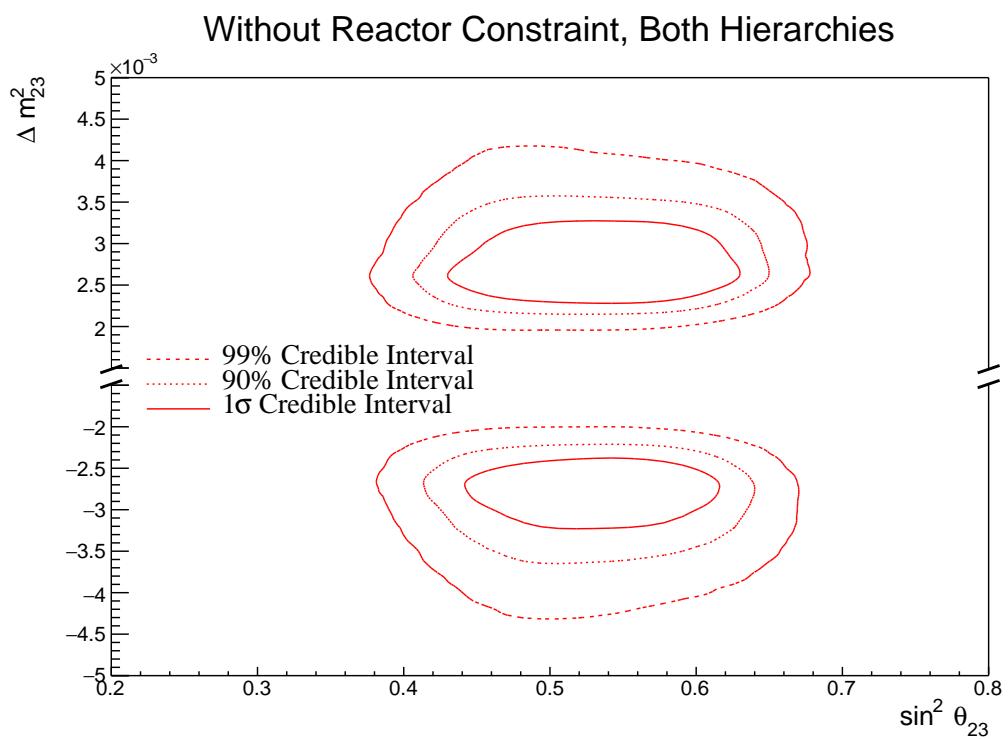


Figure 8.11: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \sin^2(\theta_{23})$, marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

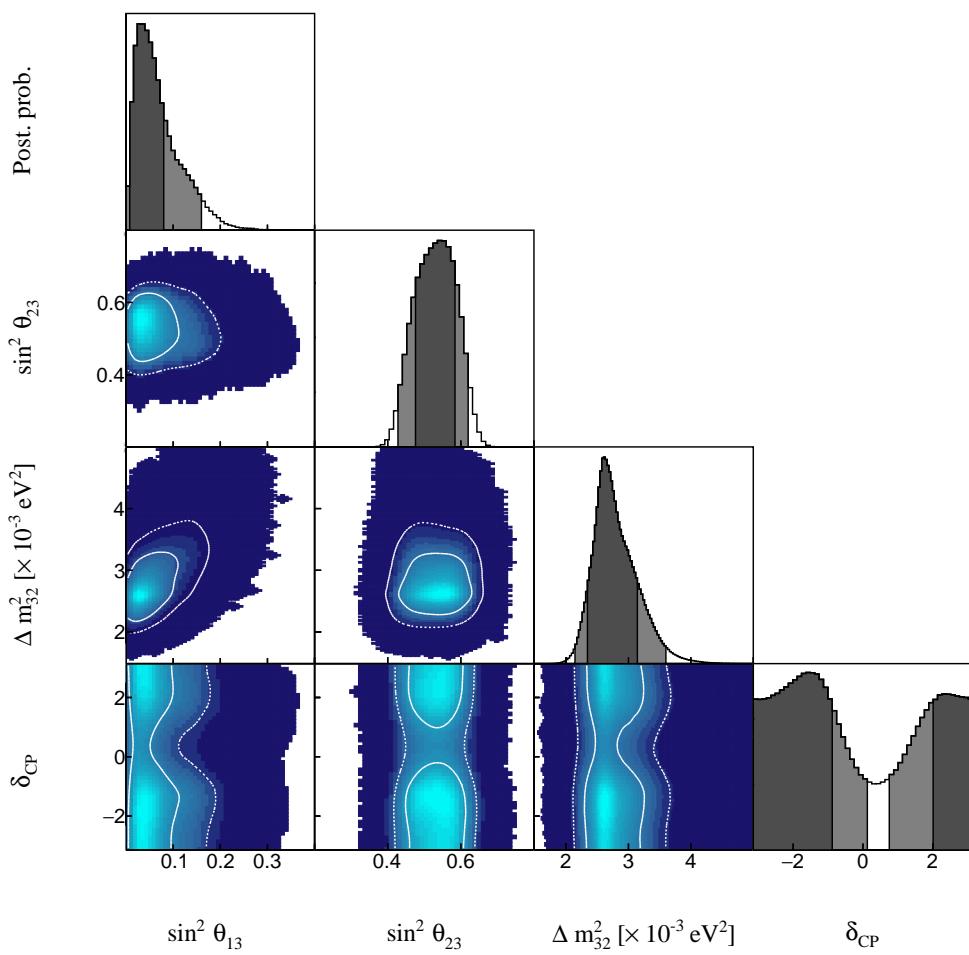


Figure 8.12: The posterior probability density distribution from the SK atmospheric-only fit. The reactor constraint is not applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

3083 8.2.2 Atmospheric-Only Sensitivity With Reactor Constraint

3084 The results in subsection 8.2.1 discuss the atmospheric sensitivity when the reactor
 3085 constraint is not applied. The correlations illustrated in Figure 8.12 indicate that
 3086 the marginalisation effects could contribute to differing sensitivities when the
 3087 external reactor constraint is applied. Using the technique discussed in subsec-
 3088 tion 4.1.1, the posterior distribution of the fit in subsection 8.2.1 can be reweighted
 3089 to include the reactor constraint of $\sin^2(\theta_{13}) = (2.18 \pm 0.08) \times 10^{-2}$ [76].

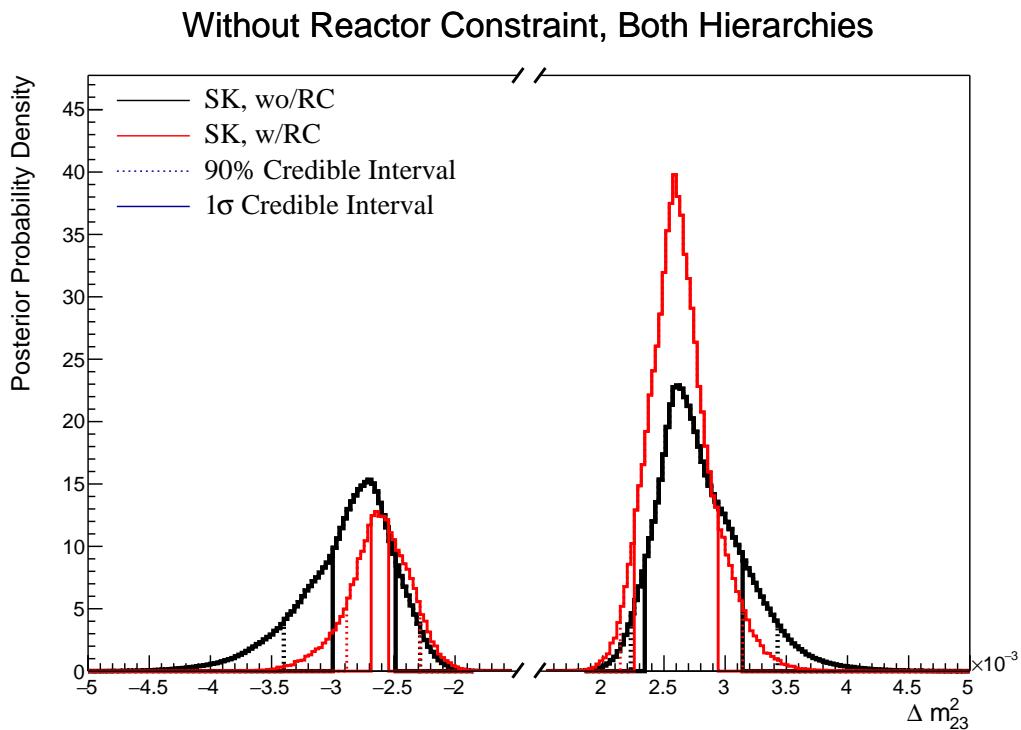


Figure 8.13: The one-dimensional posterior probability density distribution in Δm_{32}^2 compared between the SK atmospheric-only fit (Black) and the SK atmospheric fit with the reactor constraint applied (Red). The distributions are marginalised over both hierarchies.

3090 The reactor constraint increases the sensitivity of the atmospheric samples to
 3091 Δm_{32}^2 as illustrated in Figure 8.13. The 1σ credible interval in Δm_{32}^2 is determined
 3092 to be $[-2.70, -2.55] \times 10^{-3} \text{ eV}^2$ and $[2.25, 2.95] \times 10^{-3} \text{ eV}^2$. The width of the IH
 3093 credible interval is reduced by $\sim 70\%$ when the reactor constraint is applied.
 3094 Due to the marginalisation effects observed in Figure 8.12, the favoured region
 3095 of Δm_{32}^2 moves closer to zero for both hierarchies. A clear explanation of this

3096 behaviour is illustrated in Figure 8.14, which illustrates the posterior distribution
 3097 in the $\Delta m_{32}^2 - \sin^2(\theta_{13})$ parameters. The correlation between Δm_{32}^2 and $\sin^2(\theta_{13})$
 3098 is such that lower values of $\sin^2(\theta_{13})$ tend towards lower values of $|\Delta m_{32}^2|$.
 3099 This moves the posterior distribution towards the known oscillation parameter
 3100 $\Delta m_{32}^2 = 2.509 \times 10^{-3} \text{ eV}^2$.

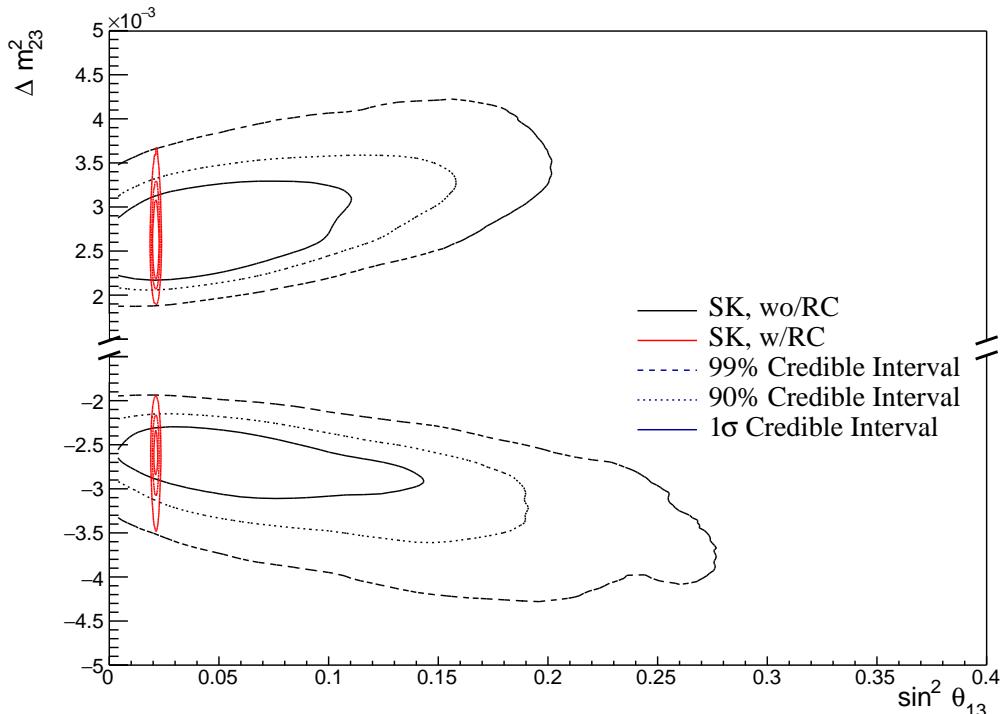


Figure 8.14: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \sin^2(\theta_{13})$ compared between the SK atmospheric-only fit (Black) and the SK atmospheric fit with the reactor constraint (Red). The distributions are marginalised over both hierarchies.

	LO ($\sin^2 \theta_{23} < 0.5$)	UO ($\sin^2 \theta_{23} > 0.5$)	Sum
NH ($\Delta m_{32}^2 > 0$)	0.21	0.53	0.74
IH ($\Delta m_{32}^2 < 0$)	0.08	0.18	0.26
Sum	0.29	0.71	1.00

Table 8.4: The distribution of steps in an SK atmospheric with reactor constraint fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The Bayes factors are calculated as $B(\text{NH}/\text{IH}) = 2.86$ and $B(\text{UO}/\text{LO}) = 2.39$.

3101 Table 8.4 presents the fraction of steps in each hierarchy and octant model
3102 for the fit after the reactor constraint has been applied. The reactor constraint
3103 significantly increases the NH preference, increasing the Bayes factor from
3104 $B(\text{NH}/\text{IH}) = 1.37$ to $B(\text{NH}/\text{IH}) = 2.86$ when the reactor constraint is applied.
3105 This is still defined as a weak preference for NH hypothesis according to Jeffrey's
3106 scale, however, it is a stronger preference than when the constraint is not applied.
3107 The preference for the correct octant model is slightly increased by the application
3108 of the reactor constraint which is consistent with expectation. However, the
3109 conclusion that would be made does not significantly change.

3110 8.2.3 Application of Near Detector Constraints for Atmospheric 3111 Samples

3112 The choice of applying the near detector constraints to the low-energy atmo-
3113 spheric samples was introduced in subsection 6.4.3. This subsection illustrates
3114 the effect of that choice on the sensitivities of the atmospheric samples to the
3115 oscillation parameters. To do this, the fit presented in subsection 8.2.1 has
3116 been compared to another fit where the constraints from the near detector
3117 have not been included. In practice, this alternative fit does not consider
3118 any near detector samples when fitting the atmospheric samples. This is the
3119 only case where the near detector constraints are neglected throughout this
3120 chapter. For both fits, the Asimov data was generated assuming the ‘AsimovA’
3121 oscillation parameter set defined in Table 2.2 and the post-BANFF systematic
3122 parameter tune.

3123 The change in sensitivity on δ_{CP} is given in Figure 8.15. The reactor constraint
3124 is not applied in either of the fits within this comparison. The shape of the
3125 posterior is similar although less peaked at the Asimov point ($\delta_{CP} = -1.601$)
3126 and more symmetric between the regions of $\delta_{CP} = -1.601$ and $\delta_{CP} \sim 2.5$. The
3127 width of the 1σ credible intervals are approximately the same (identical to within
3128 a bin width) and the same conclusion holds for the higher credible intervals. The
3129 change in sensitivity to other oscillation parameters has been studied and no
3130 significant discrepancies were found. As expected, the sensitivities are statistics
3131 dominated and the exact choice of systematic model and constraint does not
3132 significantly affect the physics conclusions one would make from this analysis.



Figure 8.15: The one-dimensional posterior probability density distribution in δ_{CP} compared between the SK atmospheric-only fit where the near detector constraint is (Black) and is not (Red) applied. The distributions are marginalised over both hierarchies.

3133 8.2.4 Atmospheric and Beam Sensitivity without Reactor Con- 3134 straint

3135 This section presents the sensitivities of the simultaneous beam and atmospheric
 3136 analysis where the reactor constraint is not applied. Similar to the previous
 3137 studies, the Asimov data is built assuming the post-BANFF cross-section tune
 3138 and Asimov A oscillation parameters defined in Table 2.2. This fit uses all 18 near
 3139 detector beam samples, 5 far detector beam samples, and 18 atmospheric samples.
 3140 The sensitivity to δ_{CP} , marginalised over both hierarchies, is given in Figure 8.16.

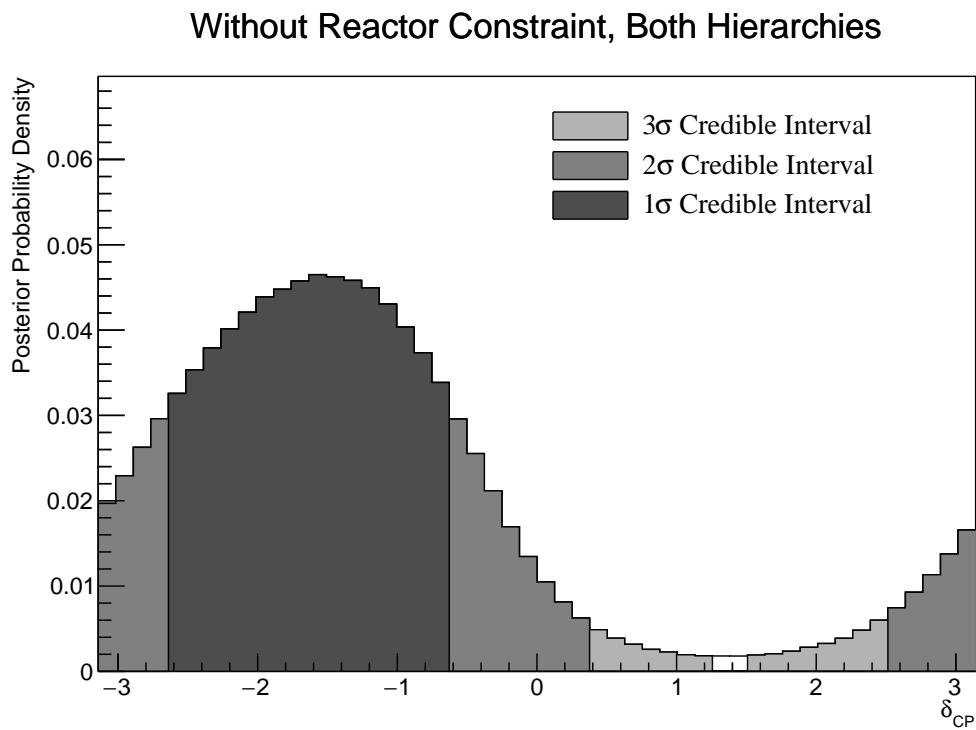


Figure 8.16: The one-dimensional posterior probability density distribution in δ_{CP} , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

3141 The credible intervals and highest posterior distribution for each oscillation
 3142 parameter is given in Table 8.5. The highest posterior probability density is
 3143 $\delta_{CP} = -1.58$ and is compatible with the known Asimov A value of $\delta_{CP} = -1.601$.
 3144 The CP-conserving values of $\delta_{CP} = 0, \pi, -\pi$ are disfavoured at 1σ credible
 3145 interval. There is also a region around $\delta_{CP} = 1.4$ which is disfavoured at more
 3146 than 3σ . Whilst these conclusions can only be made at this particular Asimov

3147 point, it does show that if the true value of δ_{CP} were CP-violating, this joint
3148 analysis would be able to disfavour CP conserving values at over 1σ without
3149 any external constraints. The highest posterior probability density does move
3150 further away from the Asimov point when only steps in the NH region are
3151 considered. This is due to the correlations between the value of δ_{CP} and the
3152 mass hierarchy, as will be later discussed.

Parameter	Interval	HPD
δ_{CP} , (BH)	$[-2.64, -0.63]$	-1.57
δ_{CP} , (NH)	$[-2.76, -0.63]$	-1.45
δ_{CP} , (IH)	$[-2.39, -0.88]$	-1.57
Δm_{32}^2 (BH) [$\times 10^{-3}\text{eV}^2$]	[2.46, 2.58]	2.49
Δm_{32}^2 (NH) [$\times 10^{-3}\text{eV}^2$]	[2.48, 2.56]	2.51
Δm_{32}^2 (IH) [$\times 10^{-3}\text{eV}^2$]	$[-2.60, -2.52]$	-2.55
$\sin^2(\theta_{23})$ (BH)	[0.48, 0.55]	0.509
$\sin^2(\theta_{23})$ (NH)	[0.48, 0.55]	0.509
$\sin^2(\theta_{23})$ (IH)	[0.48, 0.55]	0.521

Table 8.5: The position of the highest posterior probability density (HPD) and width of the 1σ credible interval for the joint beam and atmospheric fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

3153 The sensitivity to Δm_{32}^2 is illustrated in Figure 8.17, marginalised over both
3154 hierarchies. Notably, the 1σ credible interval is entirely contained within the
3155 normal hierarchy region, as illustrated in Table 8.5. This illustrates reasonable
3156 sensitivity to the mass hierarchy. This is also reflected in the 1σ credible intervals
3157 being approximately the same when they are made considering both hierarchies
3158 and when considering only the NH. The known oscillation parameter is $\Delta m_{32}^2 =$
3159 $2.509 \times 10^{-3}\text{eV}^2$. The normal hierarchy distribution favours this value with the
3160 highest posterior probability density of $\Delta m_{32}^2 = 2.51 \times 10^{-3}\text{eV}^2$.

3161 The fraction of steps in each of the mass hierarchy regions and octants of
3162 $\sin^2(\theta_{23})$ is given in Table 8.6. The Bayes factors are determined to be $B(\text{NH}/\text{IH}) =$

3163 3.67 and $B(\text{UO}/\text{LO}) = 1.74$. Jeffrey's scale states that this value of the hierarchy
 3164 Bayes factor illustrates substantial evidence for the normal hierarchy hypothesis.
 3165 This corresponds to the correct hypothesis given the known oscillation
 3166 parameters. It is a stronger statement than the atmospheric-only analysis can
 3167 provide. It is important to note that this is a substantial preference that requires
 3168 no external constraints. The Bayes factor for octant determination represents a
 3169 weak preference for the upper octant but does select the correct octant model.

	LO ($\sin^2 \theta_{23} < 0.5$)	UO ($\sin^2 \theta_{23} > 0.5$)	Sum
NH ($\Delta m_{32}^2 > 0$)	0.29	0.50	0.79
IH ($\Delta m_{32}^2 < 0$)	0.08	0.13	0.21
Sum	0.37	0.63	1.00

Table 8.6: The distribution of steps in a joint beam and atmospheric fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as $B(\text{NH}/\text{IH}) = 3.67$ and $B(\text{UO}/\text{LO}) = 1.74$.

3170 The sensitivity to $\sin^2(\theta_{23})$ is presented in Figure 8.18. There is a clear prefer-
 3171 ence for the upper octant but the peak of the distribution is relatively flat. It peaks
 3172 at $\sin^2(\theta_{23}) = 0.509$ which is in the region of the known value of $\sin^2(\theta_{23}) = 0.528$.
 3173 The difference in the highest posterior distribution and the width of the credible
 3174 interval is relatively unchanged when considering different hierarchy hypotheses
 3175 showing no strong correlation between $\sin^2(\theta_{23})$ and $|\Delta m_{32}^2|$.

3176 The sensitivity presented as a function of the appearance parameters ($\sin^2(\theta_{13}) -$
 3177 δ_{CP}) is given in Figure 8.19. As expected, the contours follow the likelihood shape
 3178 given in Figure 8.2, where the 2σ credible intervals have a closed contour exclud-
 3179 ing the region around $\delta_{CP} \sim 1.2$. The width of the 3σ credible interval in $\sin^2(\theta_{13})$
 3180 is dependent upon the value of δ_{CP} . Close to the Asimov point, $\delta_{CP} = -1.601$, the
 3181 width of the 3σ credible interval approximately spans $\sin^2(\theta_{13}) = [0.013, 0.04]$.
 3182 This is reduced to a region of $\sin^2(\theta_{13}) = [0.023, 0.042]$ at the most disfavoured
 3183 value of δ_{CP} . The 1σ credible interval is consistent with both the known oscillation

Without Reactor Constraint, Both Hierarchies

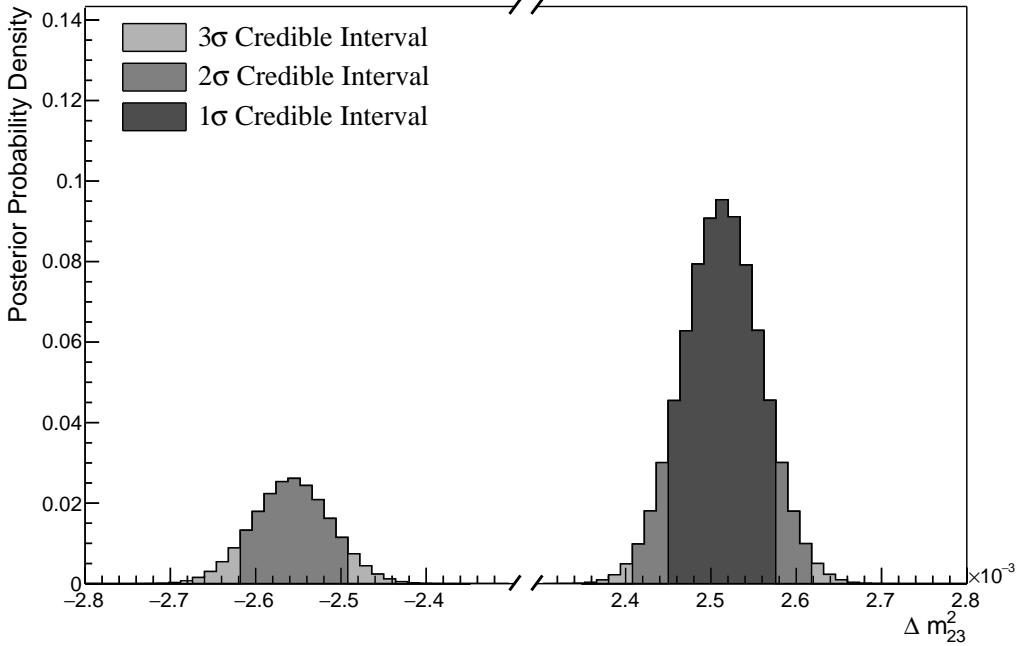


Figure 8.17: The one-dimensional posterior probability density distribution in Δm_{32}^2 , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

parameter and the reactor constraint ($\sin^2(\theta_{13}) = (2.18 \pm 0.08) \times 10^{-2}$). Application of the reactor constraint would be expected to decrease the width of the 1 σ credible intervals in δ_{CP} due to the triangular shape of the posterior probability.

The sensitivity in terms of the ‘disappearance’ parameters marginalised over both hierarchies is given in Figure 8.20. The area contained within the IH credible intervals is significantly smaller than those in the NH region. This is reflected in the IH credible intervals being tighter in the $\sin^2(\theta_{23})$ dimension. No significant correlation is observed between the value of $\sin^2(\theta_{23})$ and $|\Delta m_{32}^2|$.

The two-dimensional posterior distribution for each permutation of the oscillation parameters of interest is given in Figure 8.21. The most notable observation is that the $\sin^2(\theta_{13})$ and $\sin^2(\theta_{23})$ are anti-correlated. If the value of $\sin^2(\theta_{13})$ was constrained closer to the known oscillation parameter value, the preferred value of $\sin^2(\theta_{23})$ would increase furthering the preference for the UO. That would move the highest posterior probability closer in line with the Asimov

Without Reactor Constraint, Both Hierarchies

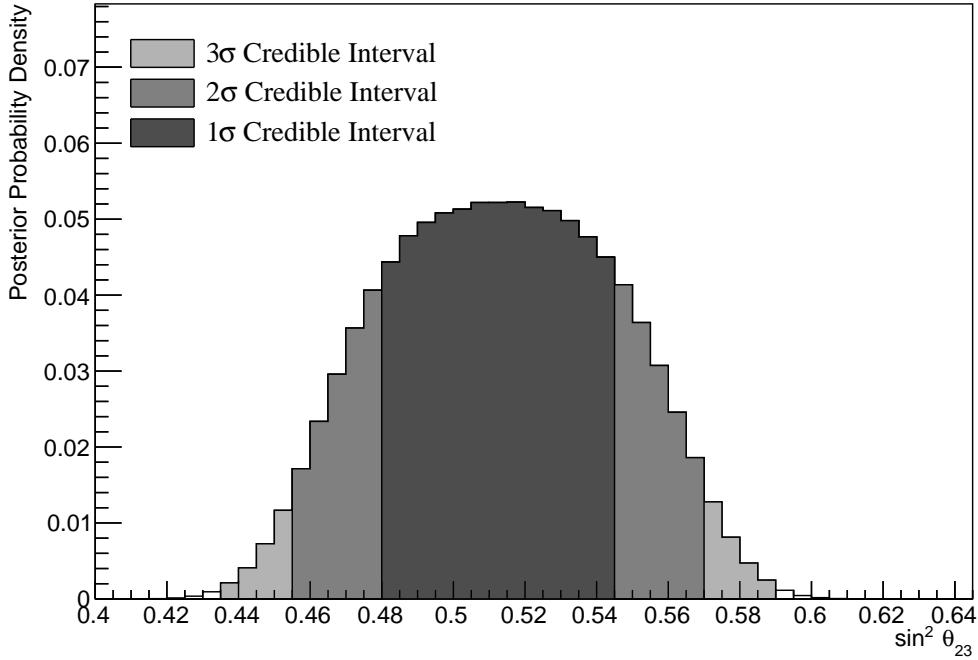


Figure 8.18: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$, marginalised over both hierarchies, from the joint beam-atmospheric fit. The reactor constraint is not applied.

3198 value. This also means that the preference for the UO would be increased if
 3199 the reactor constraint were to be applied.

3200 Furthermore, the δ_{CP} and $|\Delta m_{32}^2|$ oscillation parameters are anti-correlated,
 3201 such that higher values of $|\Delta m_{32}^2|$ prefer lower values of δ_{CP} . Whilst this is
 3202 an interesting result on its own, the width of the Δm_{32}^2 contours also depend
 3203 on $\sin^2(\theta_{13})$. This introduces another correlation effect that could modify the
 3204 sensitivity to δ_{CP} once the reactor constraint is applied.

3205 The correlation between $\sin^2(\theta_{13})$ and Δm_{32}^2 can be seen in Figure 8.22. A
 3206 much larger fraction of the posterior distribution is contained in the NH for lower
 3207 values of $\sin^2(\theta_{13})$. Consequently, the application of the reactor constraint would
 3208 be expected to significantly increase the preference for NH.

Without Reactor Constraint, Both Hierarchies

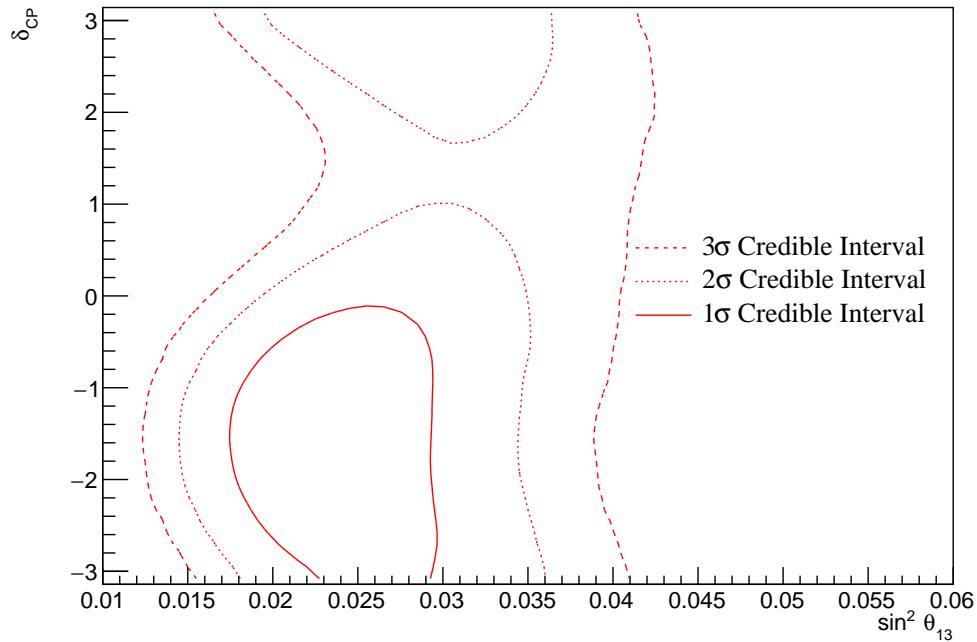


Figure 8.19: The two-dimensional posterior probability density distribution in δ_{CP} – $\sin^2(\theta_{13})$, marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

Without Reactor Constraint, Both Hierarchies

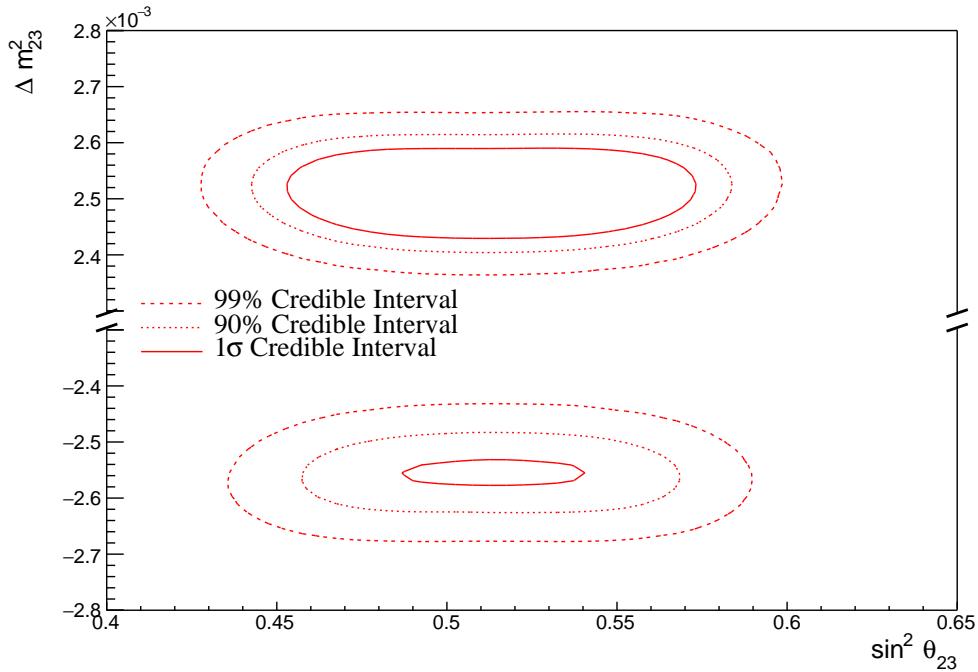


Figure 8.20: The two-dimensional posterior probability density distribution in Δm_{32}^2 – $\sin^2(\theta_{23})$, marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

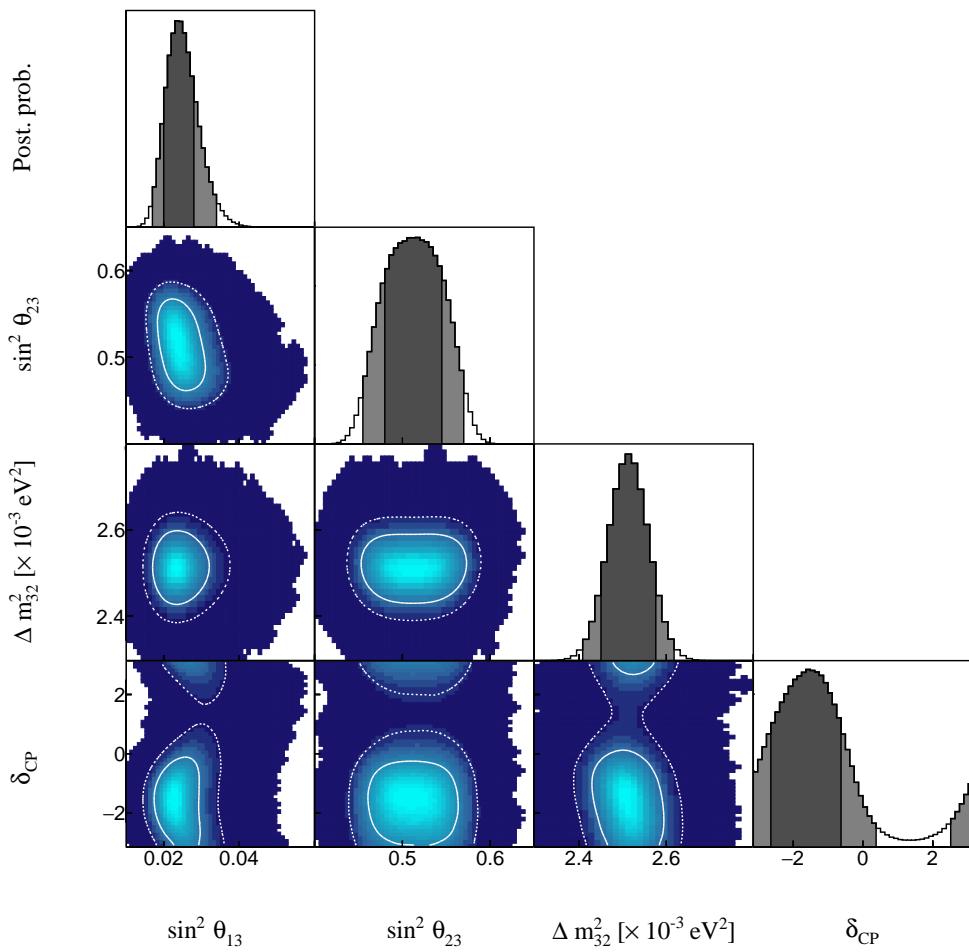


Figure 8.21: The posterior probability density distribution from the joint beam and atmospheric fit. The reactor constraint is not applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

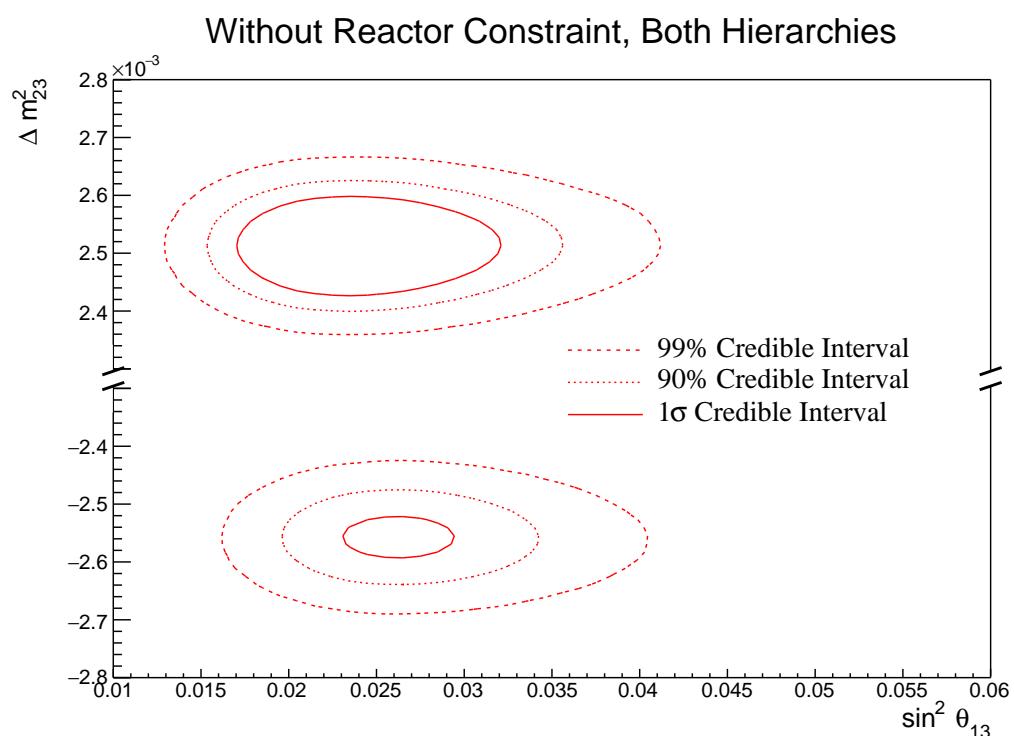


Figure 8.22: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \sin^2(\theta_{13})$, marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

3209 8.2.5 Atmospheric and Beam Sensitivity with Reactor Constraint

3210 This section presents the sensitivities of the joint beam and atmospheric fit when
 3211 the reactor constraint is applied to $\sin^2(\theta_{13})$. As with the previous studies, the
 3212 Asimov data is made using the AsimovA oscillation parameter set defined in
 3213 Table 2.2 and the post-BANFF systematic parameter tune.

With Reactor Constraint, Both Hierarchies

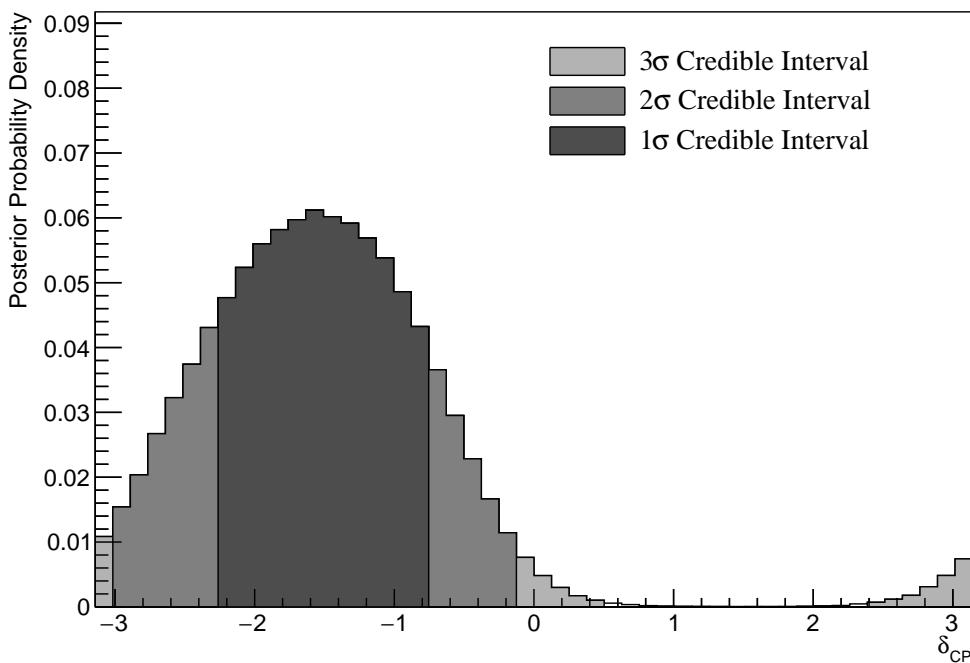


Figure 8.23: The one-dimensional posterior probability density distribution in δ_{CP} , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

3214 Figure 8.23 illustrates the sensitivity to δ_{CP} , marginalised over both hierarchies.
 3215 The CP-conserving values of $\delta_{CP} = -\pi, 0, \pi$ are disfavoured at 2σ . Furthermore,
 3216 the 3σ credible interval excludes the region of $\delta_{CP} = [0.50, 2.39]$, thus clearly
 3217 disfavouring the region of $\delta_{CP} = \pi/2$ at more than 3σ for this particular set of
 3218 known oscillation parameters. The width of the 1σ credible intervals and the
 3219 position of the highest posterior probability density is given in Table 8.7. The
 3220 highest posterior probability density in δ_{CP} is calculated as $\delta_{CP} = -1.57$ showing
 3221 no significant biases in the determination of the known oscillation parameters.

The effect of applying the reactor constraint for δ_{CP} in the joint beam-atmospheric fit is presented in Figure 8.24. The posterior distribution from the two fits are marginalised over both hierarchies. Clearly, the reactor constraint improves the ability of the fit to select the known oscillation parameter as the shape of the distribution is much more peaked around the known value of $\delta_{CP} = -1.601$. This behaviour is evidenced by the tightening of the 1σ and 90% credible intervals. Additionally, the disfavoured region of $1 < \delta_{CP} < 2$ is wider when the reactor constraint is applied. This behaviour follows from the correlations shown in Figure 8.19, where a lower value of $\sin^2(\theta_{13})$ results in tighter constraints on δ_{CP} .

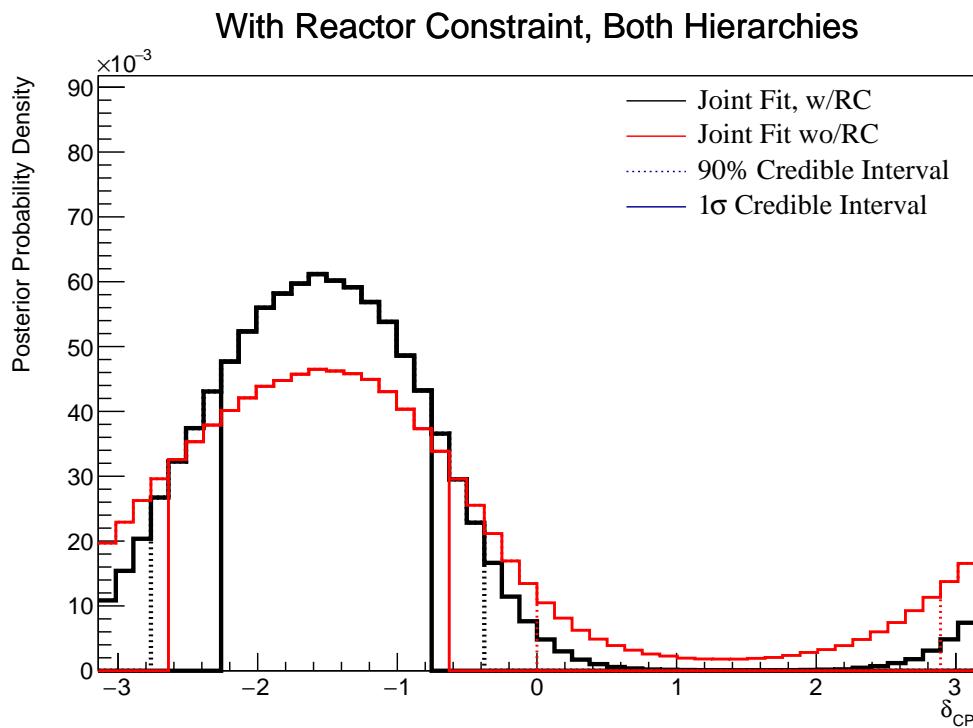


Figure 8.24: The one-dimensional posterior probability density distribution in δ_{CP} compared between the joint beam-atmospheric fit (Red) and the joint beam and atmospheric fit with the reactor constraint (Black). The distributions are marginalised over both hierarchies.

The sensitivity to $\sin^2(\theta_{23})$, marginalised over both hierarchies, is given in Figure 8.25. The highest posterior probability density is located at $\sin^2(\theta_{23}) = 0.527$ which agrees with the known value of $\sin^2(\theta_{23}) = 0.528$. The distribution clearly favours the UO with almost the entirety of the 1σ credible interval contained in the

Parameter	Interval	HPD
δ_{CP} , (BH)	[-2.26, -0.75]	-1.57
δ_{CP} , (NH)	[-2.26, -0.75]	-1.57
δ_{CP} , (IH)	[-2.13, -1.00]	-1.57
Δm_{32}^2 (BH) [$\times 10^{-3}\text{eV}^2$]	[2.46, 2.52]	2.49
Δm_{32}^2 (NH) [$\times 10^{-3}\text{eV}^2$]	[2.48, 2.56]	2.51
Δm_{32}^2 (IH) [$\times 10^{-3}\text{eV}^2$]	[-2.60, -2.52]	-2.55
$\sin^2(\theta_{23})$ (BH)	[0.49, 0.55]	0.527
$\sin^2(\theta_{23})$ (NH)	[0.49, 0.55]	0.527
$\sin^2(\theta_{23})$ (IH)	[0.50, 0.56]	0.539

Table 8.7: The position of the highest posterior probability density (HPD) and width of the 1σ credible interval for the joint beam and atmospheric fit where the reactor constraint is applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

region. Figure 8.26 highlights the sensitivity of the joint fit both with and without the reactor constraint. The fit where the reactor constraint is applied selects the known value much better. Furthermore, the reactor constraint increases the UO preference which is evidenced by the distribution moving further away from the octant boundary. This indicates that there are marginalisation effects between the two mixing parameters. This follows from the correlation illustrated between $\sin^2(\theta_{23}) - \sin^2(\theta_{13})$ in Figure 8.21. The posterior distribution of the fit with reactor constraint is more peaked compared to the flatter distribution when the reactor constraint is not applied.

The fraction of steps contained within the two hierarchy and two octant hypotheses is given in Table 8.8. The reactor constraint significantly reduces the fraction of steps that are contained within the IH-LO region from 0.08 to 0.02, whilst significantly increasing the fraction of steps within the NH-UO region from 0.53 to 0.64. The application of the reactor constraint increases the Bayes factor from $B(\text{NH}/\text{IH}) = 3.67$ to $B(\text{NH}/\text{IH}) = 7.29$. There is a very clear preference for the NH, with the Jeffreys scale stating a substantial preference for both fits.

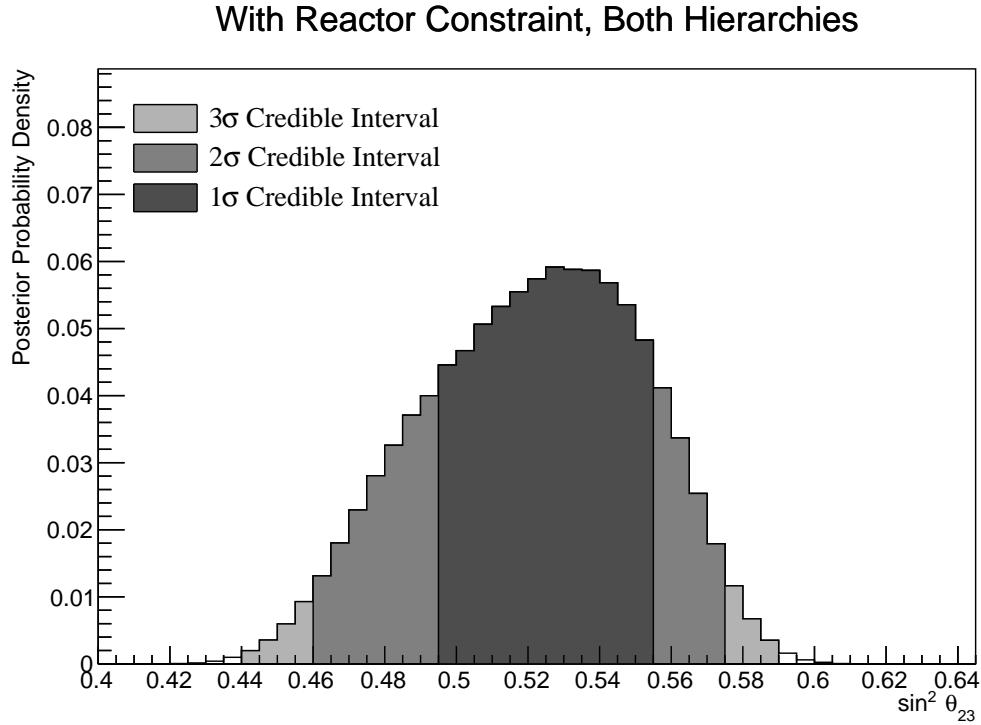


Figure 8.25: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$, marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

3251 The Bayes factor for UO preference is calculated as $B(\text{UO}/\text{LO}) = 2.86$. Whilst
 3252 still a weak preference, this is certainly a stronger statement than the sensitivity
 3253 when the reactor constraint is not applied.

	LO ($\sin^2 \theta_{23} < 0.5$)	UO ($\sin^2 \theta_{23} > 0.5$)	Sum
NH ($\Delta m_{32}^2 > 0$)	0.24	0.64	0.88
IH ($\Delta m_{32}^2 < 0$)	0.02	0.10	0.12
Sum	0.26	0.74	1.00

Table 8.8: The distribution of steps in a joint beam and atmospheric with the reactor constraint fit applied, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The Bayes factors are calculated as $B(\text{NH}/\text{IH}) = 7.29$ and $B(\text{UO}/\text{LO}) = 2.86$.

3254 The sensitivity to Δm_{32}^2 , with the reactor constraint applied, is presented in
 3255 Figure 8.27. The posterior distribution is marginalised over both hierarchies. As
 3256 expected, the 1 σ credible interval is entirely contained within the NH region. The

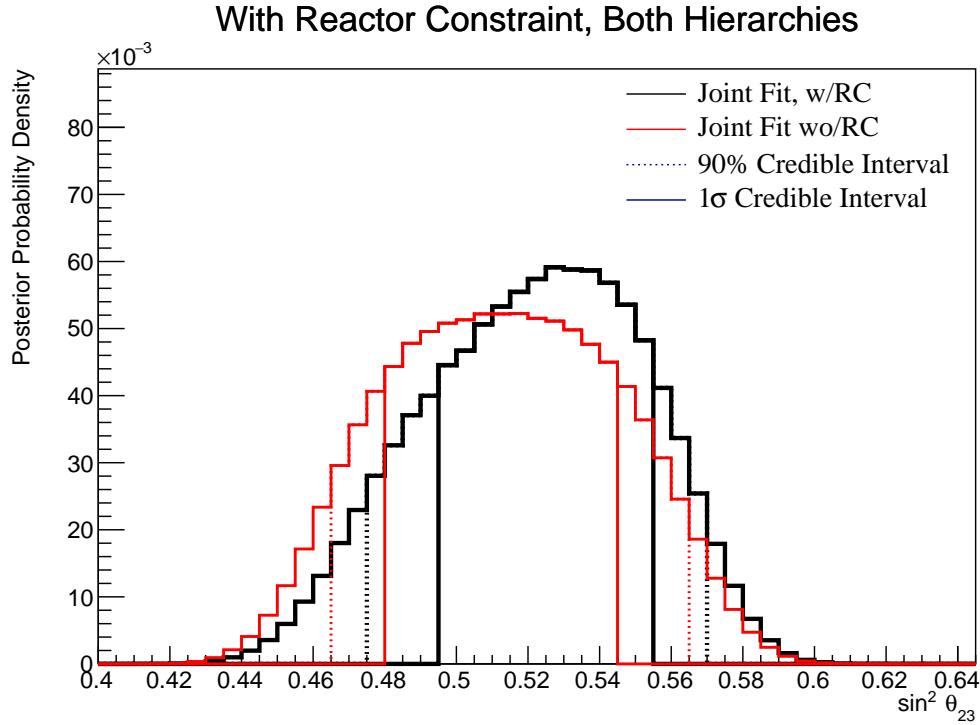


Figure 8.26: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Red) and the joint beam and atmospheric fit with the reactor constraint (Black). The distributions are marginalised over both hierarchies.

position of the highest posterior probability density is given as $2.49 \times 10^{-3} \text{ eV}^2$, illustrating no significant bias between the fit results and the known oscillation parameters. The application of the reactor constraint does not significantly move the position of the credible intervals but does reduce their width.

The sensitivity to the appearance parameters ($\sin^2(\theta_{13}) - \delta_{CP}$) is given in Figure 8.28. The distribution is mostly uncorrelated between the two parameters and is centered at the known oscillation parameters. The 1σ credible interval excludes $\delta_{CP} = 0$ and $\delta_{CP} = \pm\pi$. Furthermore, the 3σ credible intervals exclude the region of $\delta_{CP} = \pi/2$.

The sensitivity to the disappearance parameters ($\sin^2(\theta_{23}) - \Delta m_{32}^2$) is illustrated in Figure 8.29. As expected from the one-dimensional distribution, the 1σ credible interval is entirely contained within the NH region. Both the NH and IH regions favour the UO, with a visually similar preference in both hierarchies. The width of the Δm_{32}^2 1σ credible interval does not significantly depend upon

With Reactor Constraint, Both Hierarchies

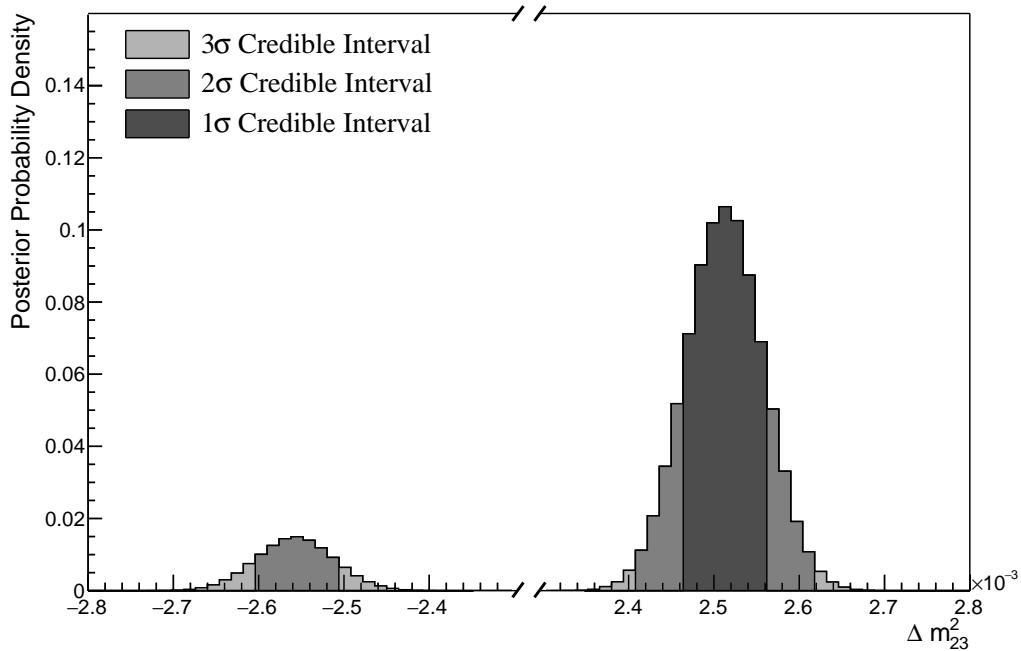


Figure 8.27: The one-dimensional posterior probability density distribution in Δm_{32}^2 , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

3271 the value or octant of $\sin^2(\theta_{23})$. This shows that there are no strong correlations
 3272 between these two parameters.

3273 Figure 8.30 illustrates the posterior distribution for each permutation of
 3274 two oscillation parameters of interest. The application of the reactor constraint
 3275 significantly reduces the correlations previously seen in Figure 8.21.

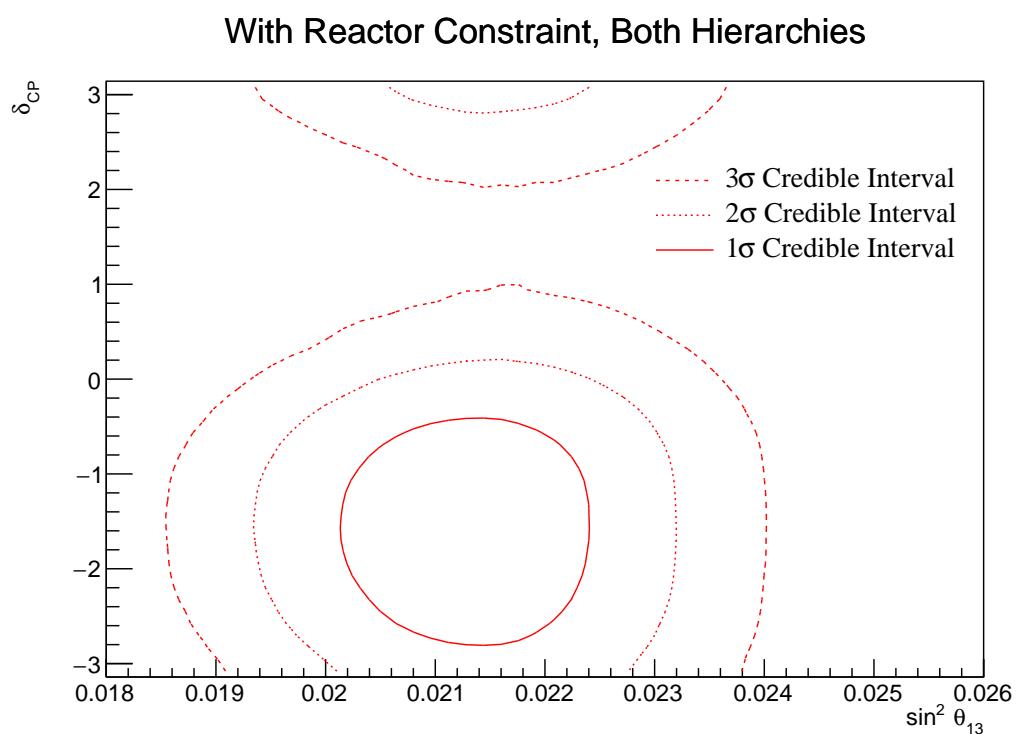


Figure 8.28: The two-dimensional posterior probability density distribution in $\delta_{CP} - \sin^2(\theta_{13})$, marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

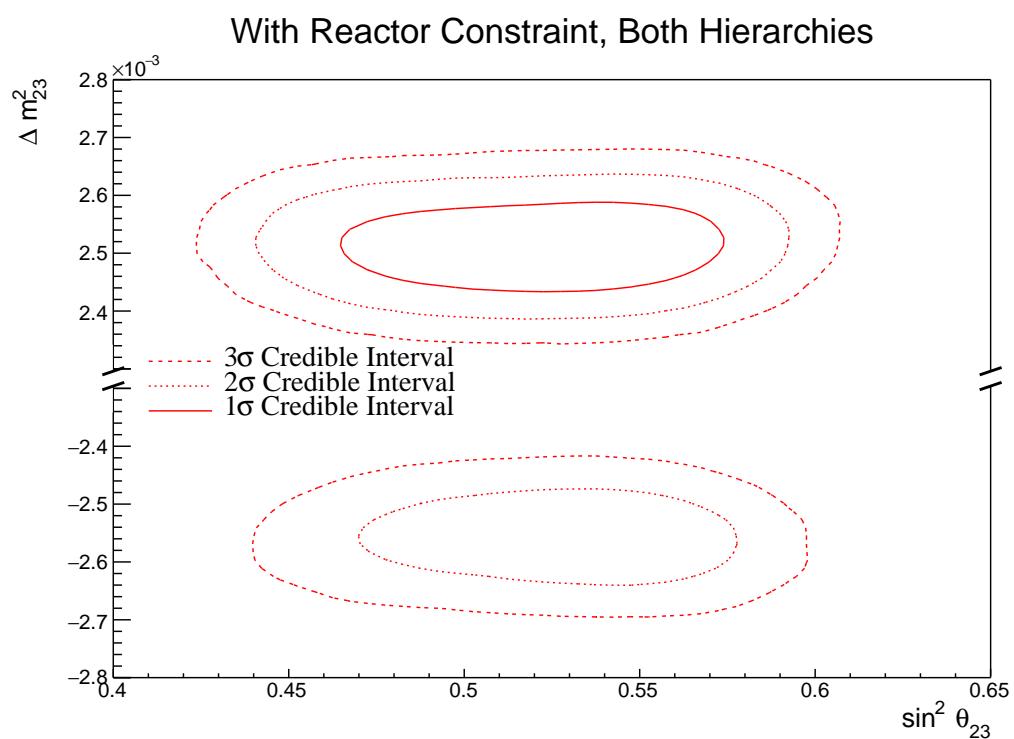


Figure 8.29: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \sin^2(\theta_{23})$, marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.



Figure 8.30: The posterior probability density distribution from the joint beam and atmospheric fit where the reactor constraint is applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

3276 8.2.6 Comparison to Latest T2K Sensitivities without Reactor 3277 Constraint

3278 The benefits of the joint beam and atmospheric analysis can be determined by
3279 comparing the sensitivities to the beam-only analysis. This section presents those
3280 comparisons for sensitivities built using the Asimov A oscillation parameters
3281 defined in Table 2.2 and the post-BANFF systematic tune. The reactor constraint
3282 is not applied within either of the fits used in these comparisons.

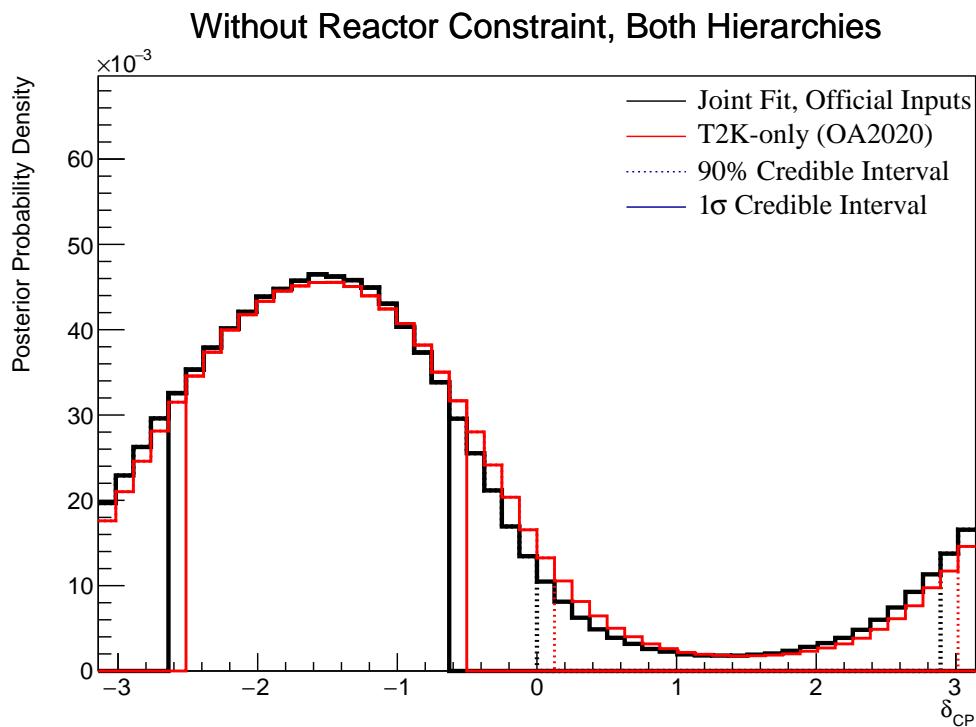


Figure 8.31: The one-dimensional posterior probability density distribution in δ_{CP} compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3283 The sensitivity, marginalised over both hierarchies, to δ_{CP} from the joint
3284 beam-atmospheric and beam-only fits is presented in Figure 8.31. As expected
3285 from the likelihood scans (Figure 8.4), the sensitivity to δ_{CP} is not significantly
3286 increased. This is because the known oscillation parameter value ($\delta_{CP} = -1.601$)
3287 lies at the position where the beam samples dominate the sensitivity compared
3288 to the SK samples.

The sensitivity to Δm_{32}^2 of the joint beam-atmospheric fit is illustrated in Figure 8.32, where the posterior distribution has been marginalised over both hierarchies. The 1σ credible interval of the joint beam and atmospheric fit is entirely contained within the NH region. This shows the significant increase in the ability of the fit to determine the correct mass hierarchy, compared to the beam-only analysis. This is further evidenced by the fact that the 90% credible intervals from the joint fit are also tighter in the IH region compared to the beam-only analysis. The Bayes factor for mass hierarchy determination for the beam-only and joint beam and atmospheric are $B(\text{NH}/\text{IH}) = 1.91$ and $B(\text{NH}/\text{IH}) = 3.67$, respectively. According to Jeffrey's scale, the beam-only analysis represents a weak preference for the NH hypothesis whereas the joint fit returns a substantial preference for the NH hypothesis. To summarise, the joint beam-atmospheric fit has a substantial preference for the correct hierarchy without the requirement of external constraints.



Figure 8.32: The one-dimensional posterior probability density distribution in Δm_{32}^2 compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3303 The sensitivity to $\sin^2(\theta_{23})$, marginalised over both hierarchies, for both the
 3304 beam-only and joint beam and atmospheric analysis are presented in Figure 8.33.
 3305 The peak of the posterior distribution from the joint analysis is more aligned with
 3306 the known value of $\sin^2(\theta_{23}) = 0.528$ compared to the beam-only analysis. The
 3307 Bayes factors for the beam-only and joint beam-atmospheric fit are $B(\text{UO/LO}) =$
 3308 1.56 and $B(\text{UO/LO}) = 1.74$, respectively. Therefore, the joint beam-atmospheric
 3309 fit does prefer the UO more strongly than the beam-only analysis, albeit slightly.

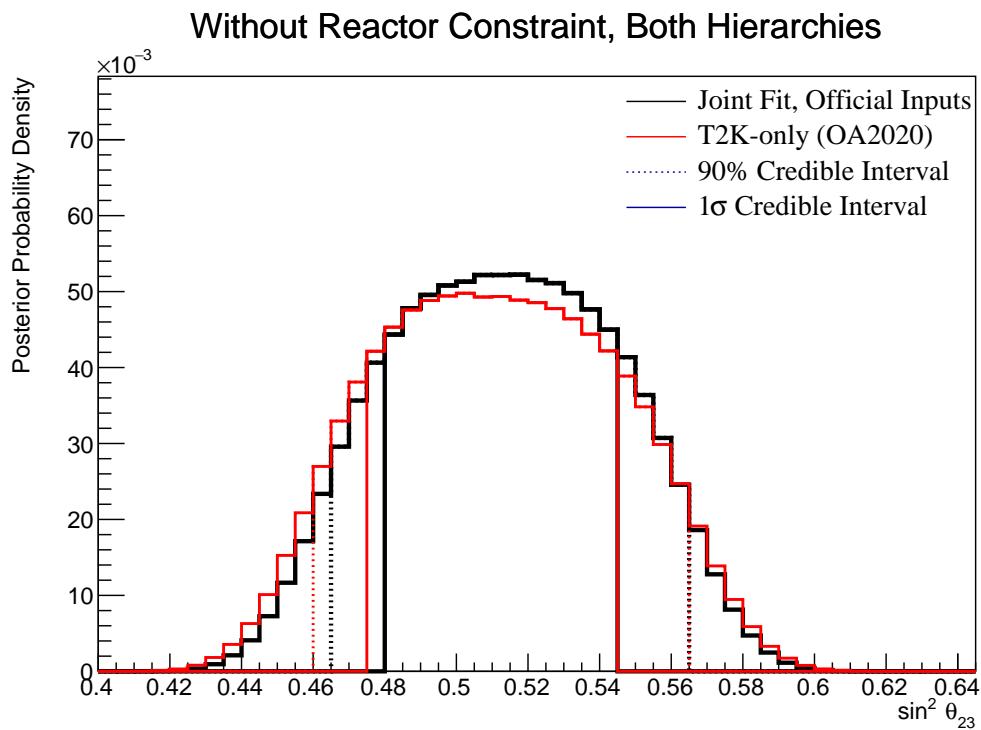


Figure 8.33: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3310 Whilst the beam-only and joint beam-atmospheric fits have similar sensitivity
 3311 to δ_{CP} and $\sin^2(\theta_{23})$ when projected in one-dimension, the benefit of the joint
 3312 analysis becomes more obvious when the sensitivities are presented in two-
 3313 dimensions. The sensitivity of the two fits to the appearance parameters ($\delta_{CP} -$
 3314 $\sin^2(\theta_{13})$) is illustrated in Figure 8.34.

3315 The width of the 99% joint fit credible interval in $\sin^2(\theta_{13})$ is squeezed in the
 3316 region of $\delta_{CP} \sim 0$ compared to the beam-only analysis. This is the same behaviour

that is seen in the appearance likelihood scans presented in Figure 8.2. The 1σ and 90% also exhibit slightly tighter constraints on δ_{CP} . This is most prevalent in the region of $\delta_{CP} \sim 0$ and $\sin^2(\theta_{13}) \sim 0.03$. Whilst the atmospheric samples do not have significant sensitivity to $\sin^2(\theta_{13})$ (as shown in Figure 8.1), they aid in breaking the degeneracy between the oscillation parameters allowing for tighter constraints.

The sensitivity to the disappearance parameters $\sin^2(\theta_{23}) - \Delta m_{32}^2$, marginalised over both hierarchies, is presented in Figure 8.35 for both the beam-only and joint beam-atmospheric fits. Whilst the one-dimensional sensitivity comparisons considered so far show the improvements of the joint fit, the two-dimensional projection really shows the benefit of adding the atmospheric samples to the beam samples. The area contained within the IH credible intervals is drastically reduced in the joint fit. This follows from the better determination of the mass hierarchy seen in the Bayes factor comparisons. The 1σ joint fit credible interval in the IH region more strongly favours the UO compared to the beam-only fit. Even in the NH region, the widths of the credible intervals in $\sin^2(\theta_{23})$ decreases, albeit to a smaller extent.

The change in sensitivity to $\delta_{CP} - \Delta m_{32}^2$ is illustrated in Figure 8.36. As expected, the contours presented within the IH region are much smaller in the joint fit due to the increased sensitivity to mass hierarchy determination. This culminates in a region around $\delta_{CP} \sim \pi/2$ which is excluded at 3σ . This behaviour is not present within the beam-only analysis. Consistent with the previous observations, the area contained within the IH credible intervals is significantly reduced in comparison to the beam-only analysis.

The sensitivity to Δm_{32}^2 and $\sin^2(\theta_{23})$, as a function of $\sin^2(\theta_{13})$, is presented in Figure 8.37 and Figure 8.38, respectively. As expected from the previous observations, the Δm_{32}^2 contours within IH region of the joint fit are much smaller than the beam-only analysis. Notably, the joint fit IH 1σ credible intervals exclude the region around the reactor constraint. This is not a bias from the fit as the known value for Δm_{32}^2 is in the NH region. This does suggest that the application

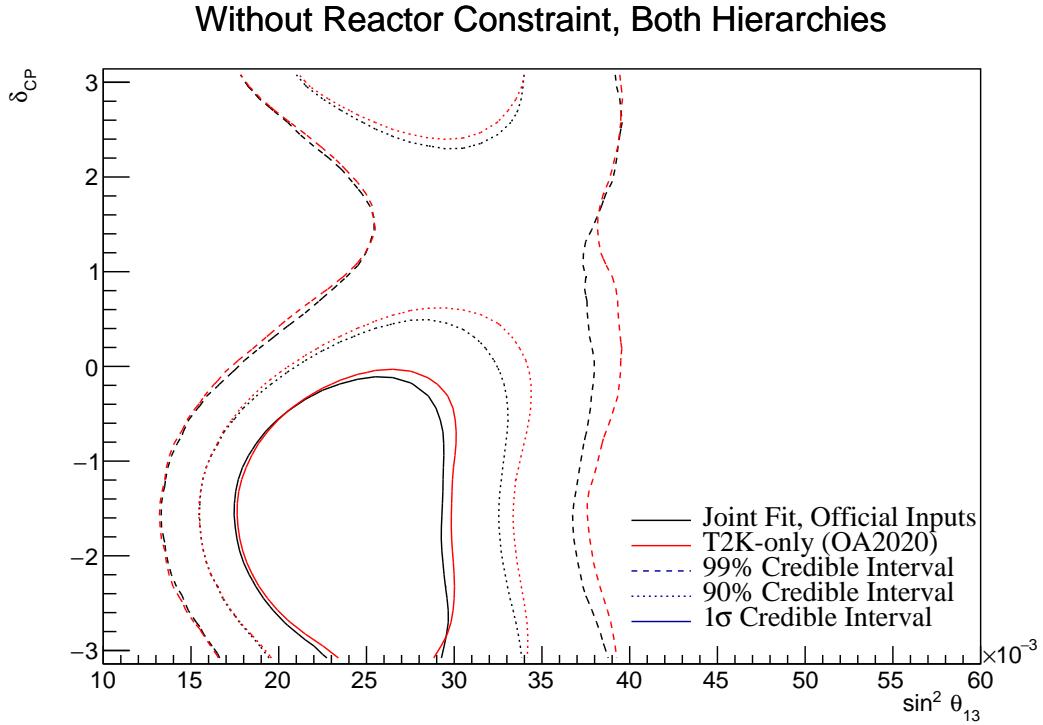


Figure 8.34: The two-dimensional posterior probability density distribution in δ_{CP} – $\sin^2(\theta_{13})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

of the reactor constraint would further increase the preference for NH in the joint fit compared to its effect on the beam-only analysis.

The beam-only and joint beam-atmospheric fits have a slightly different contour shape between the $\sin^2(\theta_{13})$ and $\sin^2(\theta_{23})$ parameters, as illustrated by Figure 8.38. The joint analysis disfavours the wrong octant hypothesis more strongly in the region of high $\sin^2(\theta_{13})$.



Figure 8.35: The two-dimensional posterior probability density distribution in Δm_{32}^2 – $\sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

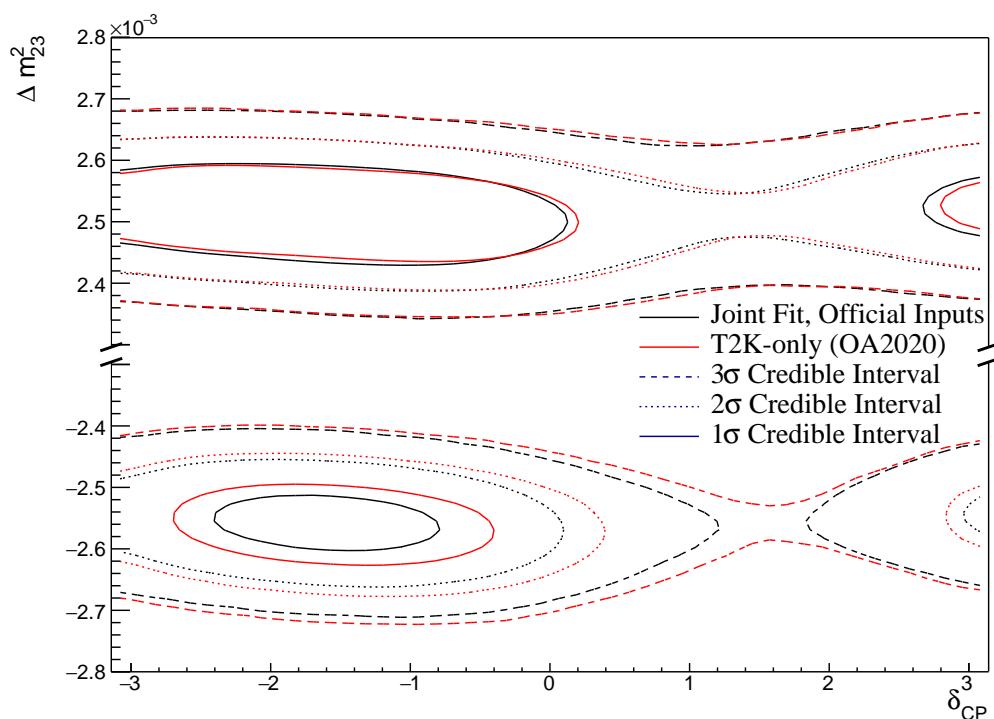


Figure 8.36: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \Delta_{CP}$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

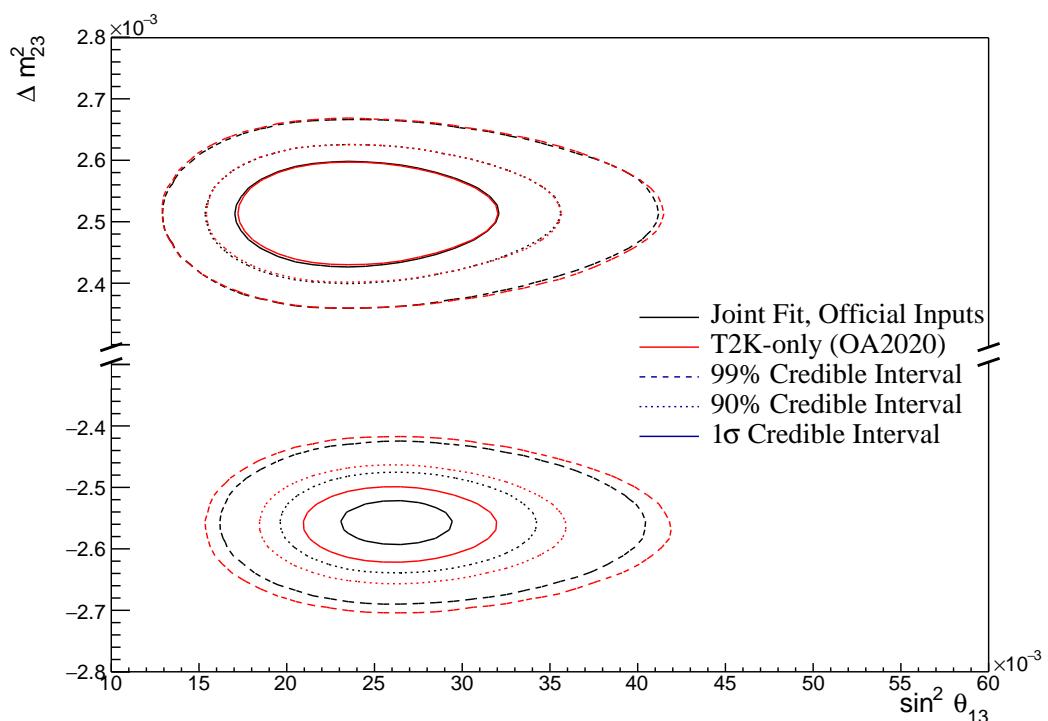


Figure 8.37: The two-dimensional posterior probability density distribution in Δm_{32}^2 – $\sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

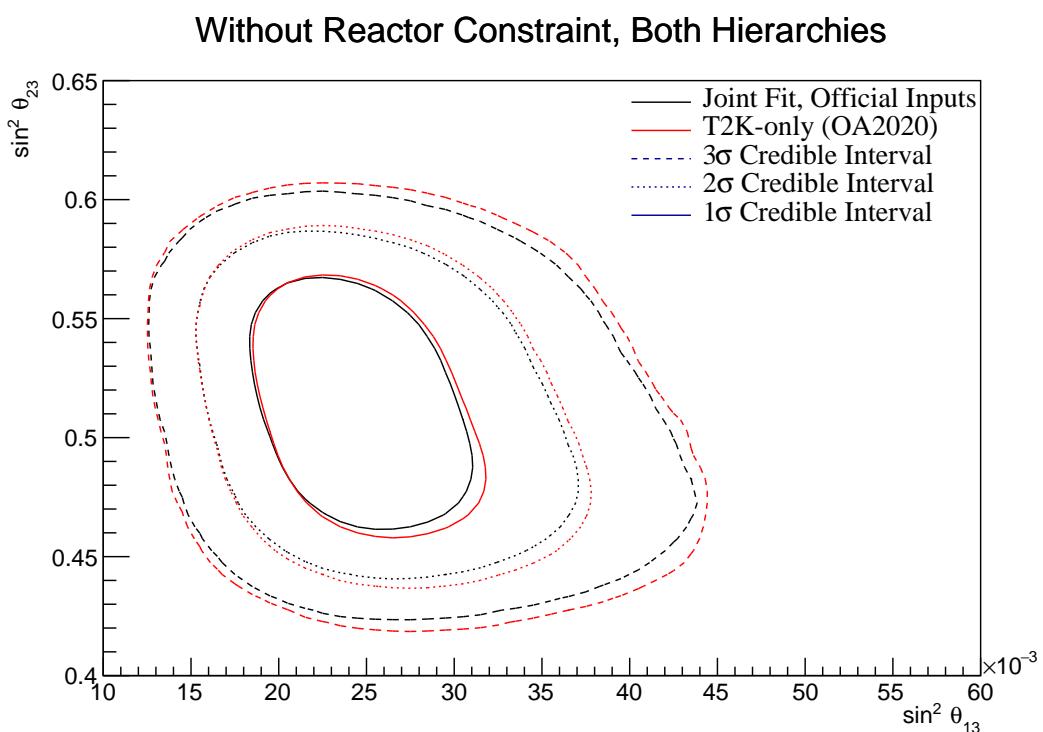


Figure 8.38: The two-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$ – $\sin^2(\theta_{13})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3353 8.2.7 Comparison to Latest T2K Sensitivities with Reactor Con- 3354 straint

3355 This section illustrates the comparison between the joint beam-atmospheric and
 3356 beam-only fits when the reactor constraint is applied. As shown in Figure 8.37,
 3357 the application of the reactor constraint is expected to significantly increase
 3358 the joint fit's preference for the NH hypothesis, compared to the beam-only
 3359 analysis. Figure 8.39 illustrates the sensitivities of the two fits to the disappearance
 3360 parameters ($\sin^2(\theta_{23}) - \Delta m_{32}^2$). This plot clearly illustrates the benefit of the joint
 3361 beam and atmospheric analysis. The 1σ credible interval in the IH region is
 3362 entirely removed in the joint analysis, illustrating the improved NH preference.

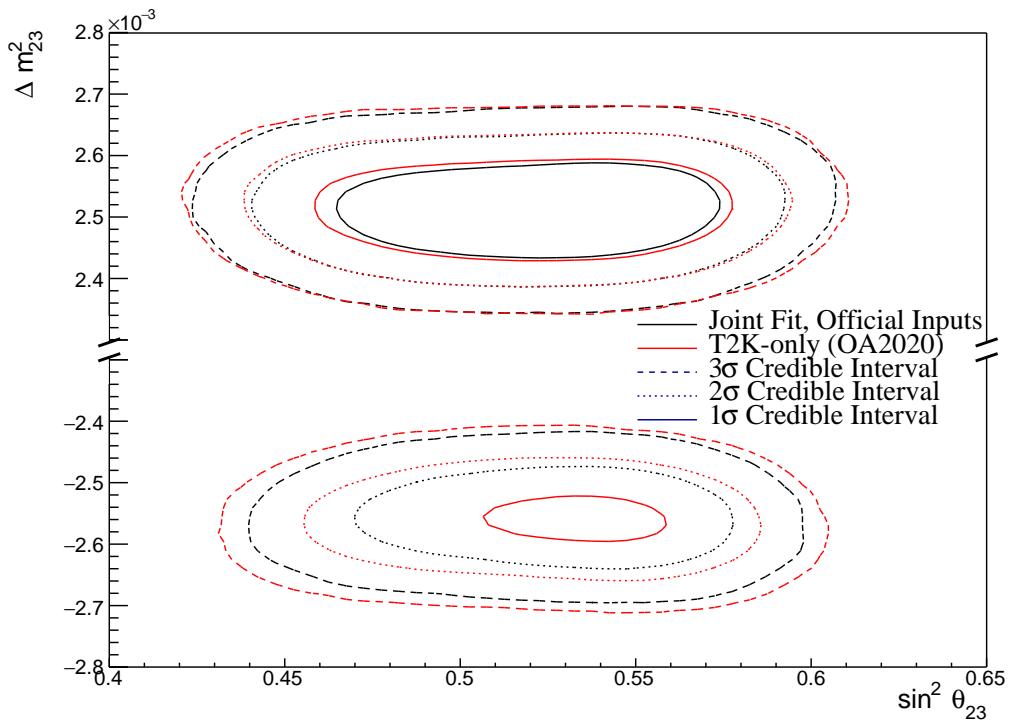


Figure 8.39: The two-dimensional posterior probability density distribution in $\Delta m_{32}^2 - \sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is applied in both fits. The distributions are marginalised over both hierarchies.

3363 The credible intervals of the joint fit are also tighter in the $\sin^2(\theta_{23})$ dimension
 3364 than the beam-only analysis in both mass hierarchy regions. This shows that
 3365 beyond the ability of the joint fit to prefer the NH more strongly than the beam-

only analysis, the precision to which it can measure $\sin^2(\theta_{23})$ is also improved. The Bayes factor for NH preference is calculated as $B(\text{NH}/\text{IH}) = 7.29$ and $B(\text{NH}/\text{IH}) = 3.41$ for the joint beam-atmospheric and beam-only analysis, respectively. Whilst both present a significant preference for the NH hypothesis, the joint fit's preference is much stronger. A similar conclusion can be made regarding the Bayes factors for UO preference which are $B(\text{UO}/\text{LO}) = 2.86$ and $B(\text{UO}/\text{LO}) = 2.67$ for the joint beam-atmospheric and beam-only analysis, respectively. Both of these represent a mild preference for the UO but there is a stronger preference observed in the joint analysis.

The sensitivity of the beam-only and joint beam-atmospheric analyses, to the appearance parameters ($\delta_{CP} - \sin^2(\theta_{13})$), are compared in Figure 8.40. These results are marginalised over both hierarchies. For this particular set of known oscillation parameters (AsimovA defined in Table 2.2), the beam-only analysis dominates the sensitivity. The joint fit does slightly increase the sensitivity to δ_{CP} but it does not change any conclusions that would be made.



Figure 8.40: The two-dimensional posterior probability density distribution in δ_{CP} – $\sin^2(\theta_{13})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is applied in both fits. The distributions are marginalised over both hierarchies.

3381 8.2.8 Alternate Asimov Parameter Set

3382 Figure 8.4 and Figure 8.5 show that the choice of the parameter set at which the
 3383 Asimov data is made can affect the conclusion. ‘AsimovA’ oscillation parameters
 3384 are defined at a region of δ_{CP} which is preferred by the T2K experiment. This
 3385 explains why the addition of the atmospheric samples does not significantly in-
 3386 crease the sensitivity to δ_{CP} , as illustrated in subsection 8.2.6 and subsection 8.2.7.
 3387 This section presents the sensitivities when ‘AsimovB’ oscillation parameters,
 3388 as defined in Table 2.2, are assumed (alongside the post-BANFF tune) when
 3389 building the Asimov data.

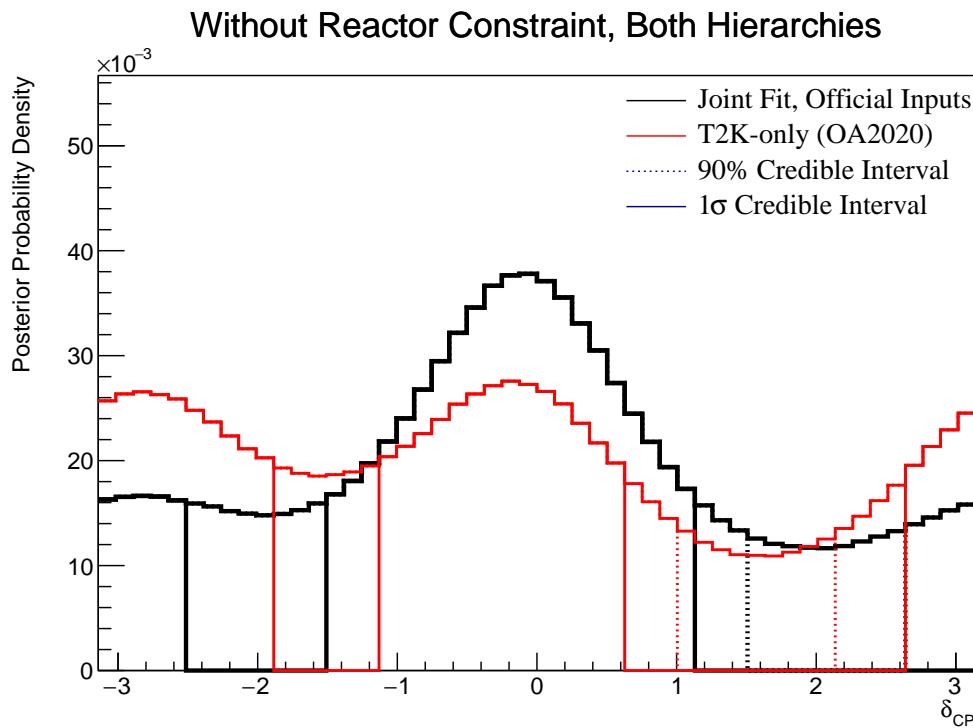


Figure 8.41: The one-dimensional posterior probability density distribution in δ_{CP} compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3390 The sensitivity to δ_{CP} for the joint beam and atmospheric fit is presented in
 3391 Figure 8.41. The results are compared to those from the beam-only analysis in
 3392 [75, 223]. The reactor constraint is not applied in either of the fits. The known
 3393 oscillation parameter value is $\delta_{CP} = 0$. The shape of the posterior distribution

from the joint analysis is more peaked at $\delta_{CP} = 0$ compared to the beam-only analysis, which has approximately the same posterior probability density at $\delta_{CP} = 0$ and $\delta_{CP} = \pi$. This shows the ability of the joint analysis to better determine the correct phase of δ_{CP} if the true value were CP-conserving. The 1σ credible intervals and the position of the highest posterior probability density are given in Table 8.9.

Parameter	Interval	HPD
δ_{CP} , (BH)	$[-\pi, -2.51], [-1.51, 1.31]$	-0.06
δ_{CP} , (NH)	$[-1.13, 1.63]$	0.06
δ_{CP} , (IH)	$[-3.02, -1.88], [-1.76, 0.13]$	-0.44
Δm_{32}^2 (BH) [$\times 10^{-3}$ eV 2]	$[-2.60, -2.49], [2.46, 2.59]$	2.51
Δm_{32}^2 (NH) [$\times 10^{-3}$ eV 2]	$[2.47, 2.56]$	2.52
Δm_{32}^2 (IH) [$\times 10^{-3}$ eV 2]	$[-2.61, -2.52]$	-2.57
$\sin^2(\theta_{23})$ (BH)	$[0.43, 0.48], [0.55, 0.59]$	0.45
$\sin^2(\theta_{23})$ (NH)	$[0.43, 0.49], [0.55, 0.58]$	0.45
$\sin^2(\theta_{23})$ (IH)	$[0.44, 0.48], [0.54, 0.59]$	0.57

Table 8.9: The position of the highest posterior probability density (HPD) and width of the 1σ credible interval for the joint beam-atmospheric fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH) and inverted hierarchy only (IH).

Naively, if just the 1σ credible interval were considered without observing the shape of the distribution, it would appear that the joint analysis would have a worse sensitivity to δ_{CP} due to the larger interval around $\delta_{CP} = 0$. The 1σ credible interval for the beam-only analysis is given as the range $\delta_{CP} = [-\pi, -1.88], [-1.13, 0.63]$ and $[2.64, \pi]$ which contains 56% of all values of δ_{CP} . The joint beam and atmospheric analysis contains 52% of all δ_{CP} values within the 1σ credible interval. Therefore, if the area within the 1σ credible interval were to be compared between the two fits, the joint analysis would be shown to have better precision.

This apparent contradiction stems from the methodology in which the credible interval is calculated. The technique used in this analysis (documented in subsection 4.3.2) fills the credible interval by selecting bins in order of probability density until 68% of the posterior density is contained. If instead, the credible interval were calculated by expanding around the highest posterior probability, the benefits of the joint fit would be more obvious. In the case where the shape of the posterior were uni-modal, these two techniques would be equivalent upto statistical fluctuations.

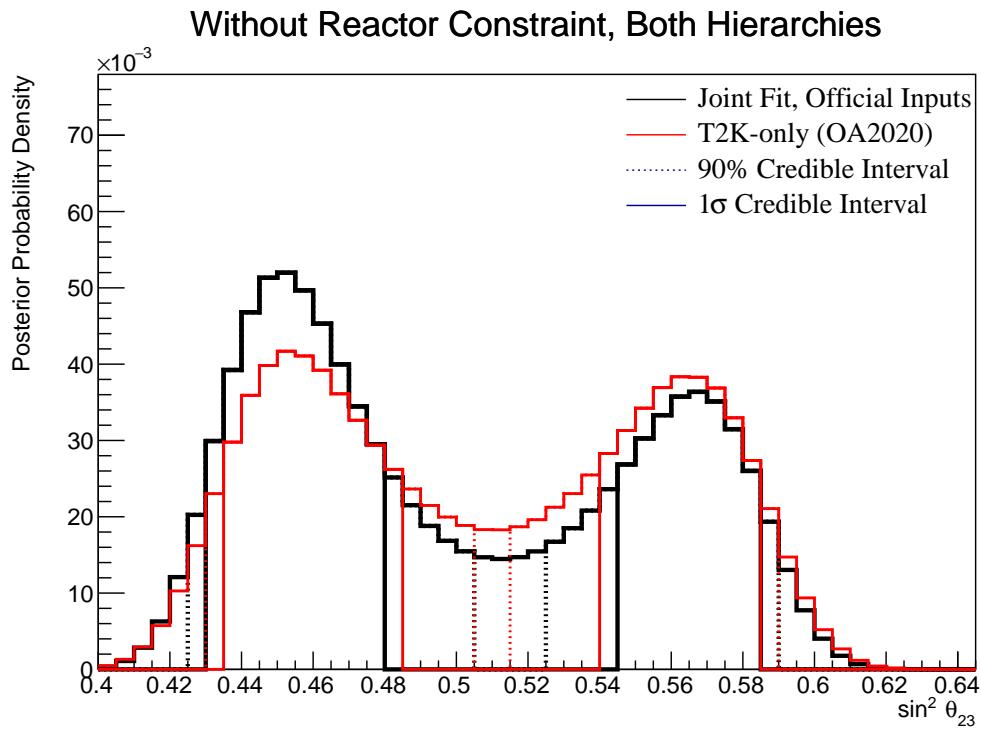


Figure 8.42: The one-dimensional posterior probability density distribution in $\sin^2(\theta_{23})$ compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

The sensitivity of the joint beam and atmospheric fit to $\sin^2(\theta_{23})$ is presented in Figure 8.42. The sensitivity is compared to that of the beam-only analysis in [75, 223]. The reactor constraint is not applied in either of the fits being compared. The Asimov parameter value is $\sin^2(\theta_{23}) = 0.45$ and the sensitivities are marginalised over both hierarchies. Clearly, the joint beam and atmospheric

3422 fit has a much larger probability density in the region surrounding the known
3423 oscillation parameter. This shows the better octant determination of the joint
3424 analysis compared to the beam-only fit. The ratio of the posterior density at the
3425 peak of the lower octant to the peak of the upper octant from the joint fit is 1.43
3426 compared to 1.09 from the beam-only analysis.

	LO ($\sin^2 \theta_{23} < 0.5$)	UO ($\sin^2 \theta_{23} > 0.5$)	Sum
NH ($\Delta m_{32}^2 > 0$)	0.35	0.24	0.59
IH ($\Delta m_{32}^2 < 0$)	0.19	0.22	0.41
Sum	0.54	0.46	1.00

Table 8.10: The distribution of steps in a joint beam and atmospheric fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as $B(\text{NH}/\text{IH}) = 1.43$ and $B(\text{LO}/\text{UO}) = 1.19$.

3427 The distribution of steps, split by hierarchy and octant hypothesis, is presented
3428 in Table 8.10. The Bayes factor for hierarchy and octant determination are
3429 $B(\text{NH}/\text{IH}) = 1.43$ and $B(\text{LO}/\text{UO}) = 1.19$, respectively. These values compare
3430 to $B(\text{NH}/\text{IH}) = 1.08$ and $B(\text{LO}/\text{UO}) = 0.91$ from the beam-only analysis. This
3431 shows additional evidence of the joint analysis's ability to select the correct octant
3432 and hierarchy hypothesis. Comparisons to the AsimovA Bayes factors presented
3433 in Table 8.6 show how the preferences for the correct octant and hierarchy depend
3434 on the true value of δ_{CP} and $\sin^2(\theta_{23})$.

3435 The sensitivity of the beam-only and joint beam-atmospheric analysis to
3436 Δm_{32}^2 is given in Figure 8.43. Both of the results are marginalised over both
3437 hierarchies and the reactor constraint is not applied in either analysis. The joint
3438 analysis has a stronger preference for the correct hierarchy (NH) which is shown
3439 by the higher Bayes factor ($B(\text{NH}/\text{IH}) = 1.43$) compared to the beam-only
3440 analysis ($B(\text{NH}/\text{IH}) = 1.08$).

Without Reactor Constraint, Both Hierarchies

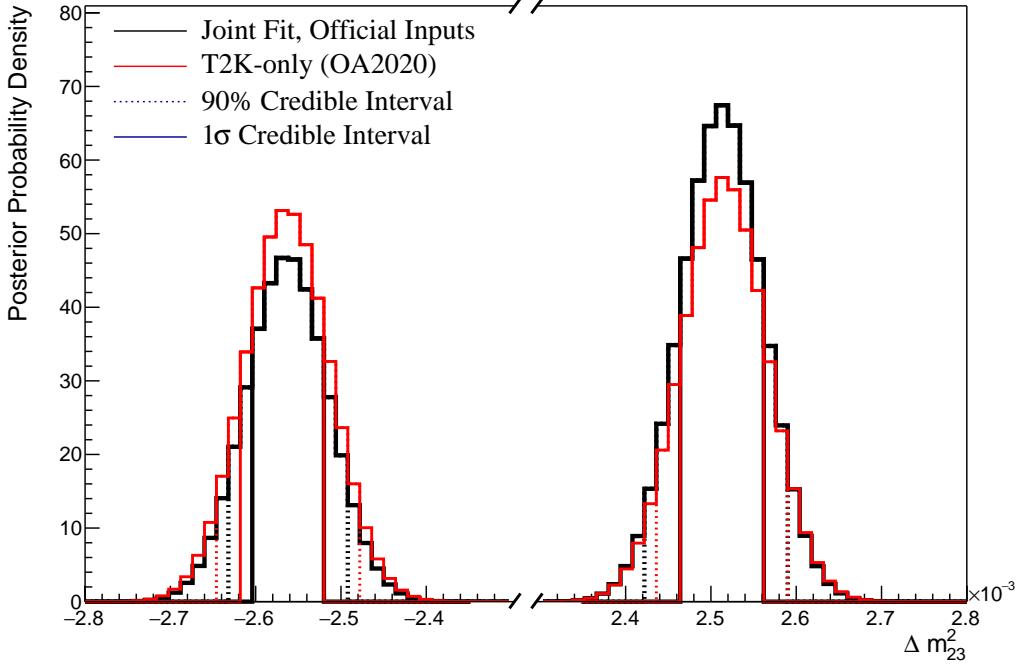


Figure 8.43: The one-dimensional posterior probability density distribution in Δm_{32}^2 compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [75, 223]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3441 8.2.9 Effect of Systematics

3442 Using the posterior predictive method documented in subsection 4.3.4, the
 3443 distribution of each sample's spectrum has been generated by sampling 2000
 3444 steps from the posterior distribution of the joint beam and atmospheric fit.
 3445 This technique reweights the Monte Carlo prediction using the systematic val-
 3446 ues given by a particular step, stores the sample spectra and repeats until the
 3447 distribution is built. The oscillation parameters are always fixed at Asimov
 3448 A values. Figure 8.44 illustrates the distribution for the SubGeV-elike-0dcy
 3449 atmospheric sample. The fit being sampled uses an Asimov data set which is
 3450 created using Asimov A oscillation parameters and the post-BANFF tune, as
 3451 detailed in subsection 8.2.4. As evidenced, the distribution does closely resemble
 3452 this Asimov data spectrum (denoted 'Post BANFF Spectra'). This would be
 3453 expected from an Asimov fit where the Monte Carlo is fit to itself but gives

³⁴⁵⁴ more credibility to the results of the fit.

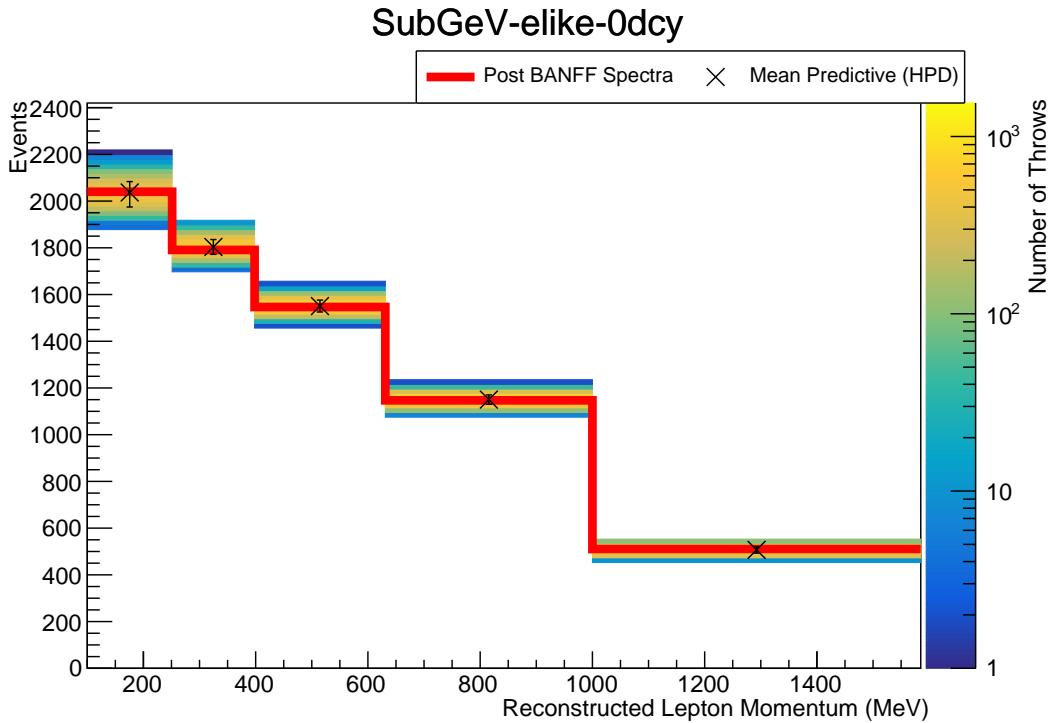


Figure 8.44: Result of the posterior predictive method for the SubGeV_elike_0dcy sample after sampling 2000 steps from the joint beam-atmospheric chain detailed in subsection 8.2.4 (Coloured histogram). The mean and uncertainty is presented for each bin. The Asimov data prediction (Red) assumes the post-BANFF tune and Asimov A oscillation parameters.

³⁴⁵⁵ The total event rate for each sample from each of the sampled steps is
³⁴⁵⁶ calculated and the fractional uncertainty, $\Delta N_i / N_i$ where N_i is the event rate
³⁴⁵⁷ of the i^{th} sample, is calculated. These values are presented in Table 8.11. In
³⁴⁵⁸ general, the impact of the systematics have an $\sim 3\%$ uncertainty on the event
³⁴⁵⁹ rate of atmospheric samples, where CC1 π targeting samples have slightly larger
³⁴⁶⁰ uncertainties than the CCQE-like samples. The fractional uncertainties on the
³⁴⁶¹ beam samples are compared to those from the beam-only analysis presented in
³⁴⁶² [75, 223]. The uncertainties on the one-ring muon samples are mostly unchanged,
³⁴⁶³ whereas the uncertainties on the one-ring electron samples are different. As
³⁴⁶⁴ discussed in subsection 8.1.1, it was discussed that the atmospheric samples
³⁴⁶⁵ would be able to add constraints on the NC background events present in the
³⁴⁶⁶ FHC1Re and RHC1Re samples. The uncertainty reduction seen in this analysis is due

³⁴⁶⁷ to that. The reason why the FHC1Re1de has a higher uncertainty in this analysis is
³⁴⁶⁸ due to the addition of the ad-hoc systematic introduce for CC1 π interactions.

Sample	Joint Analysis	Beam-only Analysis
SubGeV-elike-0dcy	2.53	-
SubGeV-elike-1dcy	3.28	-
SubGeV-mulike-0dcy	2.62	-
SubGeV-mulike-1dcy	2.23	-
SubGeV-mulike-2dcy	3.96	-
SubGeV-pi0like	2.84	-
MultiGeV-elike-nue	5.14	-
MultiGeV-elike-nuebar	2.79	-
MultiGeV-mulike	2.99	-
MultiRing-elike-nue	2.94	-
MultiRing-elike-nuebar	2.83	-
MultiRing-mulike	2.89	-
MultiRingOther-1	2.70	-
PCStop	3.22	-
PCThrus	2.99	-
UpStop-mu	2.95	-
UpThruNonShower-mu	2.70	-
UpThruShower-mu	3.19	-
FHC1Rmu	2.49	2.33
RHC1Rmu	2.89	2.93
FHC1Re	4.12	4.57
RHC1Re	5.15	5.65
FHC1Re1de	13.38	11.51

Table 8.11: The fractional uncertainty, $\Delta N / N$, as calculated from sampling 2000 throws from a joint beam-atmospheric chain. The same values for the beam samples are provided from the beam-only analysis [75, 223]. These uncertainties consider all systematic parameters to be sampled from the fit whilst the oscillation parameters are fixed at the Asimov A oscillation set.

³⁴⁶⁹ Beyond the impact on the uncertainty of each sample's event rate, the un-
³⁴⁷⁰ certainty on each systematic parameter should checked to understand how the
³⁴⁷¹ addition of atmospheric samples effects the constraints. Figure 8.45 illustrates the
³⁴⁷² central value and uncertainty on a select group of interaction systematics, for both
³⁴⁷³ the joint beam-atmospheric (from subsection 8.2.4) and the beam-only analysis.
³⁴⁷⁴ From the discussion in subsection 8.1.1, the uncertainty on systematics which are
³⁴⁷⁵ strongly constrained by the near detector should not significantly change when
³⁴⁷⁶ adding the atmospheric analysis. This behaviour is evidenced by the fact that

3477 the ratio of constraints between the two fits are very similar (within a few %) for
3478 almost all systematics. The only systematic which is more constrained in the joint
3479 beam-atmospheric analysis is the *NCOtherSK* normalisation parameter, which
3480 has a $O(10\%)$ tighter constraint. As expected, the atmospheric samples have
3481 been able to constrain this systematic which leads to the reduction in uncertainty
3482 for the beam electron-like samples.

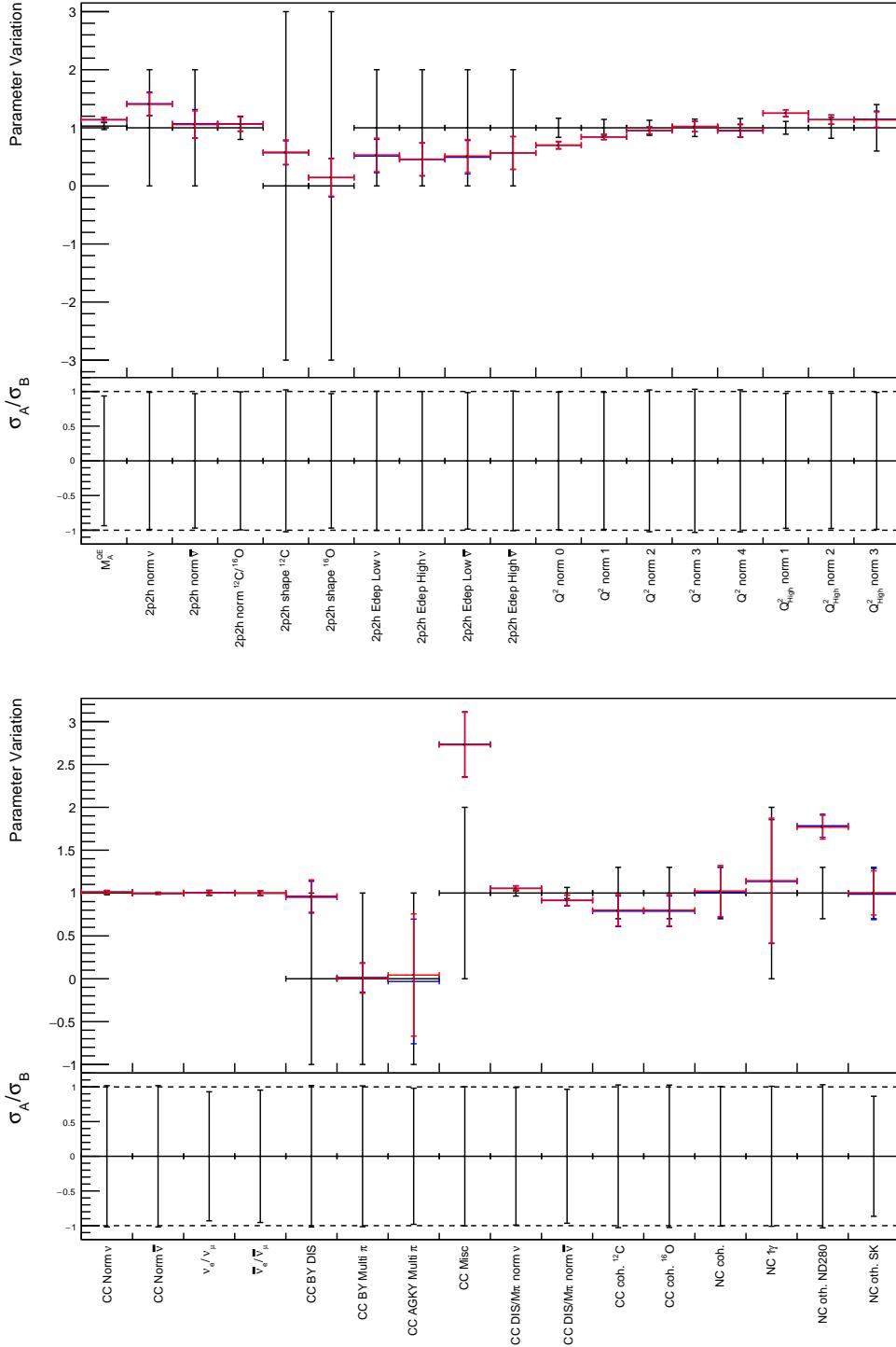


Figure 8.45: Central values and 1σ uncertainties for a select group of interaction systematics. The constraints from the prior uncertainty (Black), joint atmospheric-beam fit given in subsection 8.2.4 (Red) and beam-only analysis [75, 223] (Blue) are presented. The top pad of each plot presents the parameter variation and the bottom pad represents the ratio of the uncertainty between the joint beam-atmospheric and beam-only fits, where a value below 1.0 means the joint fit has a tighter constraint than the beam-only analysis.

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3484

Conclusions and Outlook

3485 This thesis has presented the sensitivites of the first offcial joint beam and
3486 atmospheric neutrino oscillation parameter measurements from the Tokai-to-
3487 Kamioka (T2K) and Super-Kamiokande (SK) collaborations. It combines the
3488 two independent analyses presented by the two independent collaborations DB:
3489 citations. This equates to a combined 3244.4 days equivalent of SK-IV livetime and
3490 1.97×10^{21} (1.63×10^{21}) POT in the neutrino(antineutrino) beam operating mode.

3491 The ND280 near detector is used to constrain the flux and cross-section
3492 systematics envoked within this oscillation analysis. It uses 1.15×10^{21} POT
3493 and 8.34×10^{20} POT in the neutrino and antineutrino running modes, respectively.
3494 These constraints are applied to both the beam far detector and low energy
3495 atmospheric samples through a correlated neutrino interaction model. This
3496 ensures that a consistent interaction model is used throughout the analysis.
3497 This is the first example of applying the T2K near detector constraints onto
3498 the SK atmospheric samples inside an oscillation analysis. DB: More physics
3499 developments

3500 This analysis implements a Bayesian Markov Chain Monte Carlo (MCMC)
3501 fitting technique built within the MaCh3 framework. This analysis has signifi-
3502 cantly developed the fitting framework, both in terms of technical features and
3503 performance. As a requirement for this analysis, the predominanately T2K fitting

3504 framework was required to simultaneously support and reweight alternative
3505 Monte Carlo samples. These developments include supporting new systematics,
3506 new oscillation probability calculations and previously unconsidered oscillation
3507 channels that saw the first tau events incorporated into the fitter. The devel-
3508 opments required to realise this analysis have been the building blocks of the
3509 frameworks expansion into other experiments.

3510 Due to the MCMC techniques used within the fit, a novel technique of
3511 calculating the atmospheric neutrino oscillation probabilities has been developed.
3512 This calculation uses a sub-sampling linear-averaging approach to ensure that
3513 the sensitivities being calculated are not biased due to insufficient Monte Carlo
3514 statistics in a region of fast varying probability. It illustrates a computationally
3515 feasible method of reliably calculating oscillation probabilities which can be
3516 utilised within any fitting framework. Further techniques, which consider
3517 uncertainties related to modelling the interaction height of the primary cosmic
3518 rays and the density of the Earth have also been implemented. Whilst of critical
3519 importance to this analysis, these techniques are a stand-alone technique and can
3520 be used in any other atmospheric oscillation analysis. Alongside these physics
3521 considerations, an alternative oscillation calculation engine has been interfaced
3522 with the framework to significantly reduce the resources required to perform
3523 this analysis. Current developments, based on the benefits illustrated within this
3524 analysis, are being considered within the T2K analysis.

3525 **DB: Physics results here**

3526 Whilst this analysis provides the first sensitivity measurement of a joint beam
3527 and atmospheric analysis, there are more improvements to be made. Since this
3528 analysis began, T2K has released an updated oscillation analysis with additional
3529 near and far detector samples alongside a more sophisticated interaction model.
3530 The overall change in oscillation parameter measurement observed by T2K
3531 was relatively minor DB: Bronner Nu2022 but the stronger constraints on the
3532 systematics could impact this joint analysis to a larger extent. This, or a particular
3533 focus on CCRES interaction modelling, could lead to a better understanding of

3534 the CC1 π samples from a physics-driven perspective rather than invoking the
3535 ad-hoc systematic used in this analysis. Further developments should consider
3536 the effect of correlating the beam and atmospheric flux uncertainties, where
3537 updates of the Bartol and Honda models are being made to realise this. The
3538 next goal for this analysis would be moving to a data fit. This would require
3539 performing studies which aim to understand the effect of the model choice on
3540 the oscillation parameter measurements. This tests whether there is freedom in
3541 the systematics model to allow alternative models to be fit therefore resulting
3542 in more reliable measurements.

3543 Beyond these model improvements, more data is available than what is
3544 assumed for this analysis. The T2K experiment has ran an additional period of
3545 two months, corresponding to an additional 1.78×10^{20} POT in neutrino mode.
3546 Similarly, there are several SK periods which have not been considered within this
3547 analysis. SK-I to SK-III contained approximately the same number of statistics as
3548 the SK-IV period used within this analysis but were neglected as the `fitQun`
3549 reconstruction algorithm has not validated for those periods. Furthermore,
3550 the SK-V era provides an additional 583 days of data-taking which could also
3551 be included within this analysis. This would require updating the detector
3552 systematics to reflect the changes brought about by the detector refurbishment.

3553 The T2K and SK experiments are continually developing. The near detector
3554 of the T2K experiment is currently undergoing development work to include
3555 new components. This lowers the required energy thresholds and improves
3556 vertex resolution. This may lead to stronger constraints on the flux and cross-
3557 section systematics used within the beam analysis, which could strengthen the
3558 sensitivities provided within this analysis. The SK-Gd era will also continue
3559 to accumulate statistics. Developments in the atmospheric sample selections
3560 may also benefit from the Gadolinium dopants as neutron capture will aid
3561 in neutrino/antineutrino separation leading to better mass hierarchy sensitiv-
3562 ity. This would require including interaction systematics for neutron capture

3563 of Gadolinium which has already started DB: Citation to SK neutron paper:
3564 Arxiv - 2209.08609.

3565 This analysis presents the sensitivities of the first joint beam and atmospheric
3566 analysis. This analysis and the supporting framework has the potential to become
3567 the basis of the oscillation analysis for future Hyper-Kamiokande experiment.

- 3568 • Predicted number of events at the FD
- 3569 • SK only results - w/wo RC
- 3570 • Bayes factor
- 3571 • Joint fit has significant preference for correct hierarchy without external
3572 constraints - T2K doesn't have that sensitivity
- 3573 • Summary table
- 3574 • Published at Nu2022
- 3575 • Further development of correlated detector model

Appendices

A

3577

3578

Atmospheric Sample Spectra

3579 This appendix documents the interaction mode breakdown of all the atmospheric
3580 samples used within the analysis. The generated tune of the model parameters
3581 and the Asimov A oscillation parameter set (defined in Table 2.2) are assumed.
3582 The livetime of SK-IV is taken to be 3244.4 days.

3583 A.1 Binning

3584 The lepton momentum and cosine zenith binning edges for the atmospheric
3585 samples used within this analysis are defined in Table A.1.

3586 A.2 Fully Contained Sub-GeV Samples

3587 The interaction mode breakdown of the fully contained Sub-GeV samples are
3588 shown in Figure A.1 and Figure A.2, for the samples with enriched CC0 π and
3589 CC1 π^\pm respectively.

3590 The CC0 π sample are dominated by CCQE events ($\sim 70\%$) with smaller
3591 contributions of 2p2h ($\sim 12\%$) and CC1 π ($\sim 10\%$) components. The energy peaks
3592 around 300 MeV, which is slightly below that of the T2K samples but still has
3593 significant contribution upto 1 GeV which overlaps the T2K sample energy range.

Sample	$\cos(\theta_Z)$ Bins	Momentum Bin Edges ($\log_{10}(P)$ MeV)
SubGeV-elike-0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-elike-1dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-1dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-2dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-pi0like	1	2.0, 2.2, 2.4, 2.6, 2.8, 3.2
MultiGeV-elike-nue	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV-elike-nuebar	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV-mulike	10	3.0, 3.4, 5.0
MultiRing-elike-nue	10	3.0, 3.4, 3.7, 5.0
MultiRing-elike-nuebar	10	3.0, 3.4, 3.7, 5.0
MultiRing-mulike	10	2.0, 3.124, 3.4, 3.7, 5.0
MultiRing-Other1	10	3.0, 3.4, 3.7, 4.0, 5.0
PC-Stop	10	2.0, 3.4, 5.0
PC-Through	10	2.0, 3.124, 3.4, 3.7, 5.0
Upmu-Stop	10	3.2, 3.4, 3.7, 8.0
Upmu-Through-Showering	10	2.0, 8.0
Upmu-Through-NonShowering	10	2.0, 8.0

Table A.1: The reconstructed cosine zenith and lepton momentum binning assigned to the atmospheric samples. The “ $\cos(\theta_Z)$ Bins” column illustrates the number of bins uniformly distributed over the $-1.0 \leq \cos(\theta_Z) \leq 1.0$ region for fully and partially contained samples and $-1.0 \leq \cos(\theta_Z) \leq 0.0$ region for up- μ samples.

3594 The one-ring CC1 π samples, where the pion is tagged via its decay electron,
 3595 are dominated by CC1 π events ($\sim 75\%$) with a small contribution of CCM π
 3596 ($\sim 10\%$). The two-ring pion sample is mostly dominated by the NC1 π^0 via
 3597 resonances, and has several equally-sized contributions from CCQE, NC1 π^\pm via
 3598 resonances, and NC coherent pion production, where the π^0 likely comes from
 3599 nucleon and π^\pm final state interactions in the nucleus.

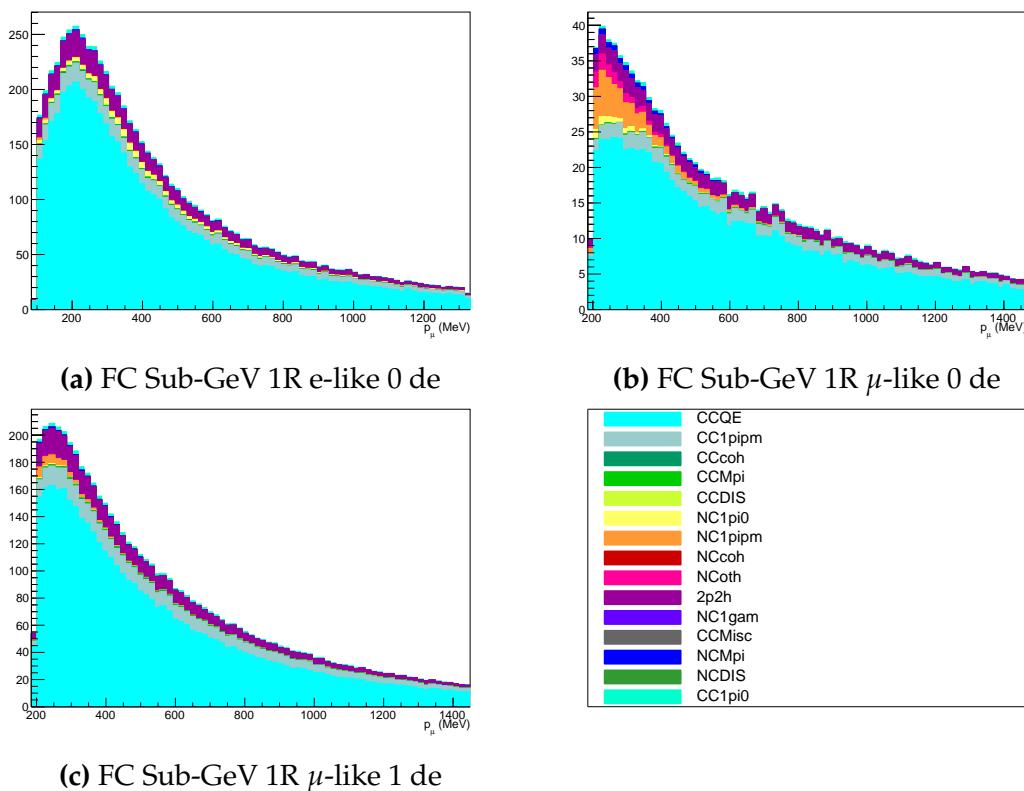


Figure A.1: Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting CC 0π events.

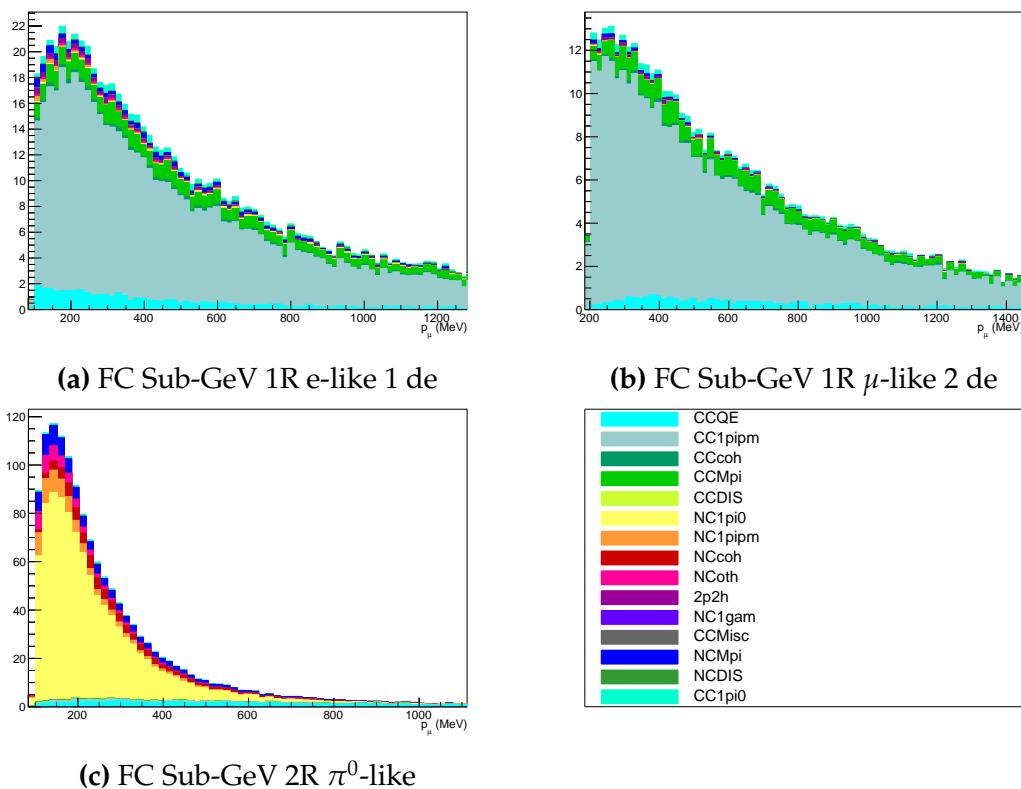


Figure A.2: Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting single pion events.

A.3 Fully Contained Multi-GeV Samples

The interaction mode breakdown of fully contained multi-GeV samples is highlighted in Figure A.3. Due to the event selection applied in SK which targets π^+ and π^- separation, the ν_e sample mainly consists of events with pions (single pion production or multi-pion/DIS interactions). The pion separation is explained in Section section 6.1. This reasoning also explains the significant CCQE contribution of the $\bar{\nu}_e$ sample. The muon-like sample is dominated by CCQE interactions with $\sim 10 - 15\%$ 2p2h and CC1 π contribution of events.

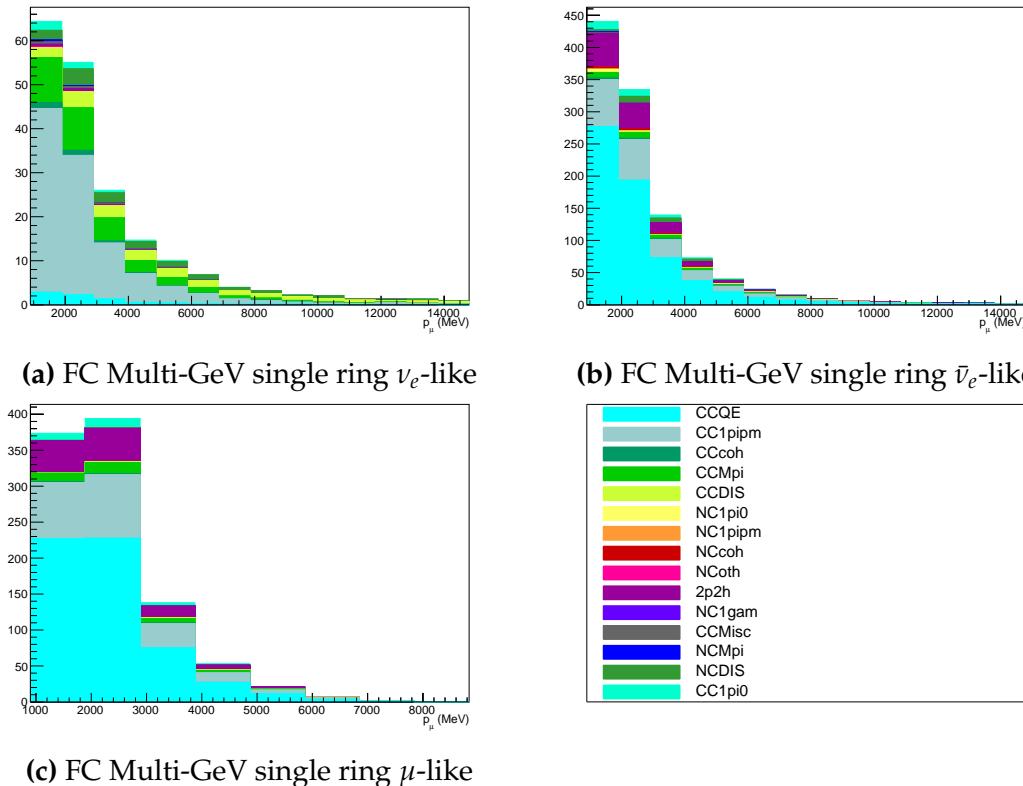


Figure A.3: Breakdown by interaction mode of the FC Multi-GeV single ring atmospheric samples.

3608 A.4 Fully Contained Multi-Ring Samples

3609 The interaction mode breakdown of fully contained multi-ring events is shown
 3610 in Figure A.4. These samples see more interaction modes contributing in general,
 3611 and there is a much larger contribution from multi-pion and DIS interaction
 3612 modes, compared to the other samples.

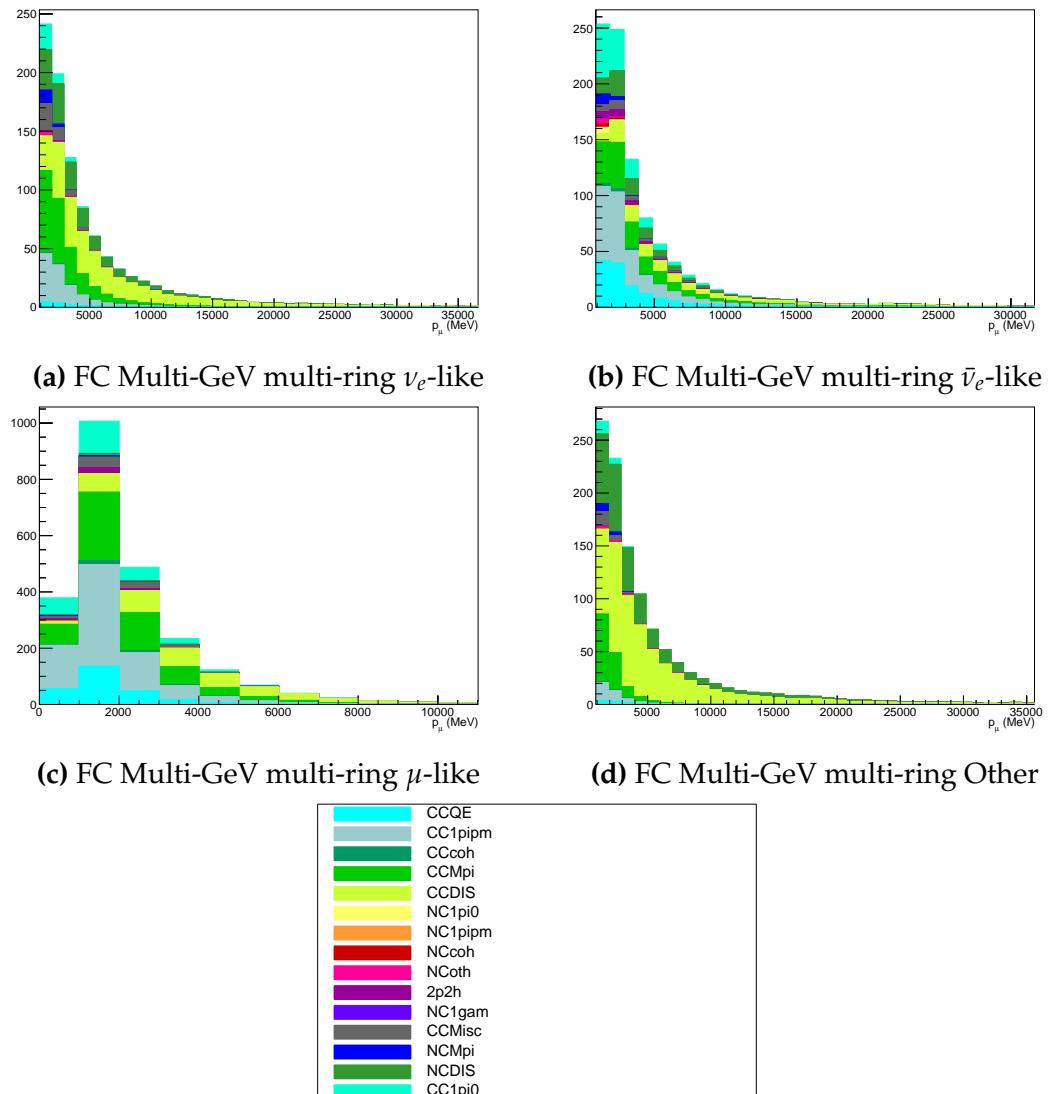


Figure A.4: Breakdown by interaction mode of the FC Multi-GeV multi-ring atmospheric samples.

³⁶¹³ A.5 Partially Contained Samples

³⁶¹⁴ The breakdown for partially contained samples is highlighted in Figure A.5.
³⁶¹⁵ As with the multi-ring samples, there is no dominating interaction mode. The
³⁶¹⁶ neutrino energies of events in this sample extend into the tens of GeV and become
³⁶¹⁷ dominated by DIS interaction modes in the high energy limit.

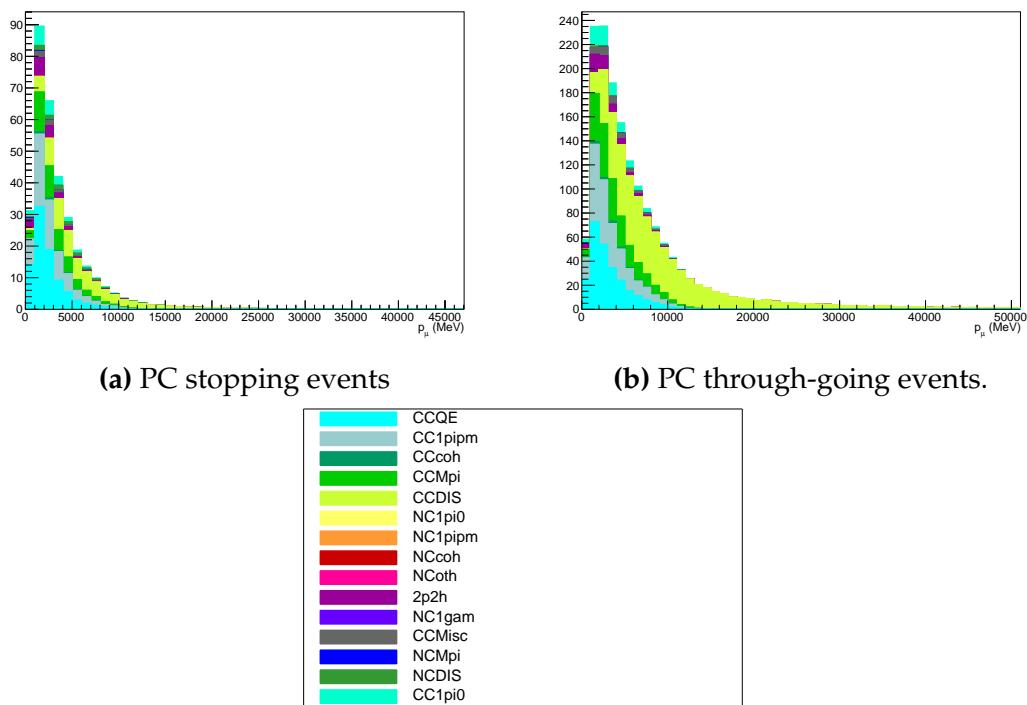


Figure A.5: Breakdown by interaction mode of the PC atmospheric samples.

3618 A.6 Upward-Going Muon Samples

3619 The breakdown for upward-going muons is illustrated in Figure A.6. These
 3620 samples are significantly dominated by DIS interactions with energies extending
 3621 up into the hundreds of GeV.

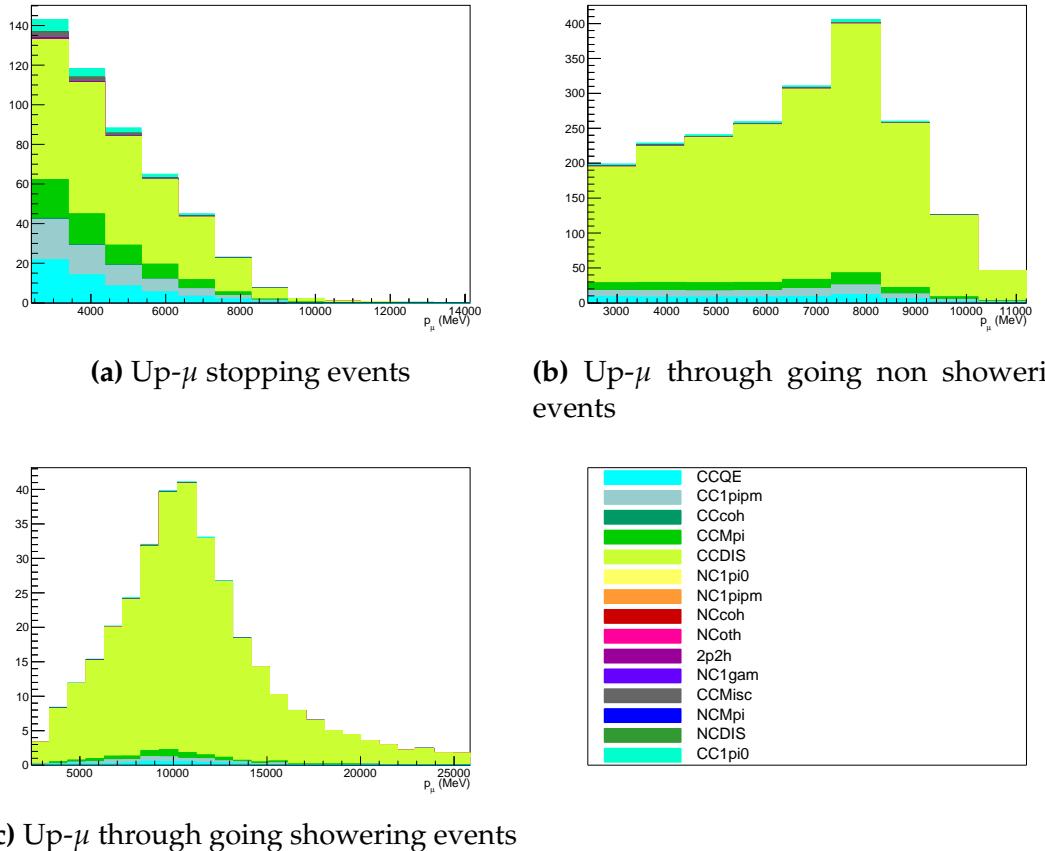


Figure A.6: Breakdown by interaction mode of the atmospheric upward going muon samples.

Bibliography

- [1] J Chadwick. "Intensitätsverteilung im magnetischen Spectrum der β -Strahlen von radium B + C". In: *Verhandl. Dtsc. Phys. Ges.* 16 (1914), p. 383. URL: <http://cds.cern.ch/record/262756>.
- [2] C D Ellis and W A Wooster. "The average energy of disintegration of radium E". en. In: *Proc. R. Soc. Lond. A Math. Phys. Sci.* 117.776 (Dec. 1927), pp. 109–123.
- [3] W. Pauli. "Dear radioactive ladies and gentlemen". In: *Phys. Today* 31N9 (1978), p. 27.
- [4] E. Fermi. "An attempt of a theory of beta radiation. 1." In: *Z. Phys.* 88 (1934), pp. 161–177.
- [5] F. Reines and C. L. Cowan. "Detection of the Free Neutrino". In: *Phys. Rev.* 92 (3 1953), pp. 830–831. URL: <https://link.aps.org/doi/10.1103/PhysRev.92.830>.
- [6] C. L. Cowan et al. "Detection of the Free Neutrino: a Confirmation". In: *Science* 124.3212 (1956), pp. 103–104. eprint: <http://science.sciencemag.org/content/124/3212/103.full.pdf>. URL: <http://science.sciencemag.org/content/124/3212/103>.
- [7] G. Danby et al. "Observation of High-Energy Neutrino Reactions and the Existence of Two Kinds of Neutrinos". In: *Phys. Rev. Lett.* 9 (1 1962), pp. 36–44. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.9.36>.
- [8] K. Kodama et al. "Observation of tau neutrino interactions". In: *Physics Letters B* 504.3 (2001), pp. 218 –224. URL: <http://www.sciencedirect.com/science/article/pii/S0370269301003070>.
- [9] A. Aguilar-Arevalo et al. "Evidence for neutrino oscillations from the observation of anti-neutrino(electron) appearance in a anti-neutrino(muon) beam". In: *Phys. Rev.* D64 (2001), p. 112007. arXiv: [hep-ex/0104049](https://arxiv.org/abs/hep-ex/0104049) [hep-ex].
- [10] A. A. Aguilar-Arevalo et al. "Improved Search for $\bar{\nu}_\mu \rightarrow \bar{\nu}_e$ Oscillations in the MiniBooNE Experiment". In: *Phys. Rev. Lett.* 110 (16 2013), p. 161801. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.110.161801>.
- [11] Planck Collaboration. "Planck 2018 results. VI. Cosmological parameters". In: *aap* 641 (Sept. 2020).
- [12] J. A. Bagger et al. "Precision electroweak measurements on the Z resonance". In: *Physics Reports* 427.5 (2006), pp. 257 –454. URL: <http://www.sciencedirect.com/science/article/pii/S0370157305005119>.
- [13] B. Pontecorvo. "Neutrino Experiments and the Problem of Conservation of Leptonic Charge". In: *Sov. Phys. JETP* 26 (1968). [Zh. Eksp. Teor. Fiz. 53, 1717(1967)], pp. 984–988.

- 3659 [14] B. Pontecorvo. "Inverse beta processes and nonconservation of lepton charge". In:
3660 *Sov. Phys. JETP* 7 (1958). [Zh. Eksp. Teor. Fiz.34,247(1957)], pp. 172–173.
- 3661 [15] M. Kobayashi and T. Maskawa. "CP-Violation in the Renormalizable Theory of
3662 Weak Interaction". In: *Progress of Theoretical Physics* 49.2 (1973), pp. 652–657. URL:
3663 <http://dx.doi.org/10.1143/PTP.49.652>.
- 3664 [16] N. Cabibbo. "Unitary Symmetry and Leptonic Decays". In: *Phys. Rev. Lett.* 10 (12
3665 1963), pp. 531–533. URL:
3666 <https://link.aps.org/doi/10.1103/PhysRevLett.10.531>.
- 3667 [17] A Maio et al. "Search for Majorana neutrinos with the SNO+ detector at
3668 SNOLAB". In: *Journal of Physics: Conference Series* 587 (2015), p. 012030. URL:
3669 <https://doi.org/10.1088/1742-6596/587/1/012030>.
- 3670 [18] A Yu Smirnov. "The MSW effect and solar neutrinos". In: (2003).
- 3671 [19] S.P. Mikheyev and A.Y. Smirnov. "Resonance enhancement of oscillations in
3672 matter and solar neutrino spectroscopy". In: *Soviet Journal of Nuclear Physics* 42 (6
3673 1985), pp. 913–917.
- 3674 [20] L. Wolfenstein. "Neutrino oscillations in matter". In: *Phys. Rev. D* 17 (9 1978),
3675 pp. 2369–2374. URL: <https://link.aps.org/doi/10.1103/PhysRevD.17.2369>.
- 3676 [21] V. D. Barger et al. "Matter Effects on Three-Neutrino Oscillations". In: *Phys. Rev.*
3677 *D* 22 (1980), p. 2718.
- 3678 [22] Y. Ashie et al. "Evidence for an Oscillatory Signature in Atmospheric Neutrino
3679 Oscillations". In: *Phys. Rev. Lett.* 93 (10 2004), p. 101801. URL:
3680 <https://link.aps.org/doi/10.1103/PhysRevLett.93.101801>.
- 3681 [23] Q. R. Ahmad et al. "Direct Evidence for Neutrino Flavor Transformation from
3682 Neutral-Current Interactions in the Sudbury Neutrino Observatory". In: *Phys.*
3683 *Rev. Lett.* 89 (1 2002), p. 011301. URL:
3684 <https://link.aps.org/doi/10.1103/PhysRevLett.89.011301>.
- 3685 [24] 2015 Nobel prize in Physics as listed by Nobelprize.org.
3686 https://www.nobelprize.org/nobel_prizes/physics/laureates/2015/.
3687 Accessed: 22-06-2022.
- 3688 [25] J. A. Formaggio and G. P. Zeller. "From eV to EeV: Neutrino Cross Sections
3689 Across Energy Scales". In: *Rev. Mod. Phys.* 84 (2012), pp. 1307–1341. arXiv:
3690 1305.7513 [hep-ex].
- 3691 [26] A Oralbaev, M Skorokhvatov, and O Titov. "The inverse beta decay: a study of
3692 cross section". In: *Journal of Physics: Conference Series* 675.1 (2016), p. 012003. URL:
3693 <https://doi.org/10.1088/1742-6596/675/1/012003>.
- 3694 [27] A Bellerive. "Review of solar neutrino experiments". en. In: *Int. J. Mod. Phys. A*
3695 19.08 (Mar. 2004), pp. 1167–1179.
- 3696 [28] R Davis, D S. Harmer, and K C. Hoffman. "Search for Neutrinos from the Sun".
3697 In: *Phys. Rev. Lett.* 20 (21 1968), pp. 1205–1209. URL:
3698 <https://link.aps.org/doi/10.1103/PhysRevLett.20.1205>.
- 3699 [29] N Vinyoles et al. "A new generation of standard solar models". In: *Astrophys. J.*
3700 835.2 (Jan. 2017), p. 202.

- 3701 [30] V Gribov and B Pontecorvo. "Neutrino astronomy and lepton charge". en. In: *Phys. Lett. B* 28.7 (Jan. 1969), pp. 493–496.
- 3702 [31] K. S. Hirata et al. "Observation of ${}^8\text{B}$ solar neutrinos in the Kamiokande-II detector". In: *Phys. Rev. Lett.* 63 (1 1989), pp. 16–19. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.63.16>.
- 3703 [32] W Hampel et al. "GALLEX solar neutrino observations: results for GALLEX IV". en. In: *Phys. Lett. B* 447.1-2 (Feb. 1999), pp. 127–133.
- 3704 [33] J. N. Abdurashitov et al. "Measurement of the solar neutrino capture rate with gallium metal". In: *Phys. Rev. C* 60 (5 1999), p. 055801. URL: <https://link.aps.org/doi/10.1103/PhysRevC.60.055801>.
- 3705 [34] Q R Ahmad et al. "Direct evidence for neutrino flavor transformation from neutral-current interactions in the Sudbury neutrino observatory". en. In: *Phys. Rev. Lett.* 89.1 (June 2002).
- 3706 [35] Borexino Collaboration. "Comprehensive measurement of pp-chain solar neutrinos". en. In: *Nature* 562.7728 (Oct. 2018), pp. 505–510.
- 3707 [36] B Aharmim et al. "A search for neutrinos from the SolarhepReaction and the diffuse supernova neutrino background with the Sudbury neutrino observatory". en. In: *Astrophys. J.* 653.2 (Dec. 2006), pp. 1545–1551.
- 3708 [37] M Agostini et al. "Experimental evidence of neutrinos produced in the CNO fusion cycle in the Sun". In: (2020).
- 3709 [38] G Danby et al. "Observation of high-energy neutrino reactions and the existence of two kinds of neutrinos". In: *Phys. Rev. Lett.* 9.1 (July 1962), pp. 36–44.
- 3710 [39] K. Abe et al. "T2K neutrino flux prediction". In: *Physical Review D* 87.1 (2013). URL: <https://doi.org/10.1103%2Fphysrevd.87.012001>.
- 3711 [40] D. G. Michael et al. "Observation of Muon Neutrino Disappearance with the MINOS Detectors in the NuMI Neutrino Beam". In: *Phys. Rev. Lett.* 97 (19 2006), p. 191801. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.97.191801>.
- 3712 [41] G. Danby et al. "Observation of High-Energy Neutrino Reactions and the Existence of Two Kinds of Neutrinos". In: *Phys. Rev. Lett.* 9 (1 1962), pp. 36–44. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.9.36>.
- 3713 [42] M. A. Acero et al. "First measurement of neutrino oscillation parameters using neutrinos and antineutrinos by NOvA". In: *Phys. Rev. Lett.* 123 (15 2019), p. 151803. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.123.151803>.
- 3714 [43] K. Abe et al. "The T2K experiment". In: *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment* 659.1 (2011), pp. 106 –135. eprint: <http://www.sciencedirect.com/science/article/pii/S0168900211011910>.
- 3715 [44] B Abi et al. "Long-baseline neutrino oscillation physics potential of the DUNE experiment". en. In: *Eur. Phys. J. C Part. Fields* 80.10 (Oct. 2020).
- 3716 [45] Hyper-Kamiokande Proto-Collaboration. "Physics potential of a long-baseline neutrino oscillation experiment using a J-PARC neutrino beam and Hyper-Kamiokande". In: *Prog. Theor. Exp. Phys.* 2015.5 (May 2015), pp. 53C02–0.

- 3745 [46] C. Blanco, D. Hooper, and P. Machado. "Constraining sterile neutrino
3746 interpretations of the LSND and MiniBooNE anomalies with coherent neutrino
3747 scattering experiments". In: *Physical Review D* 101.7 (2020). URL:
3748 <https://doi.org/10.1103%2Fphysrevd.101.075051>.
- 3749 [47] MicroBooNE Collaboration. *Search for an Excess of Electron Neutrino Interactions in*
3750 *MicroBooNE Using Multiple Final State Topologies*. 2021. URL:
3751 <https://arxiv.org/abs/2110.14054>.
- 3752 [48] B. Armbruster et al. "Upper limits for neutrino oscillations $\bar{\nu}_\mu \rightarrow \bar{\nu}_e$ from muon
3753 decay at rest". In: *Phys. Rev. D* 65 (11 2002), p. 112001. URL:
3754 <https://link.aps.org/doi/10.1103/PhysRevD.65.112001>.
- 3755 [49] T. K. Gaisser and M. Honda. "Flux of Atmospheric Neutrinos". In: (2002).
- 3756 [50] G. D. Barr et al. "Three-dimensional calculation of atmospheric neutrinos". In:
3757 *Physical Review D* 70.2 (2004). URL:
3758 <https://doi.org/10.1103/physrevd.70.023006>.
- 3759 [51] M. Honda et al. "Calculation of atmospheric neutrino flux using the interaction
3760 model calibrated with atmospheric muon data". In: *Physical Review D* 75.4 (2007).
3761 URL: <https://doi.org/10.1103/physrevd.75.043006>.
- 3762 [52] M. Honda et al. "New calculation of the atmospheric neutrino flux in a
3763 three-dimensional scheme". In: *Phys. Rev. D* 70 (4 2004), p. 043008. URL:
3764 <https://link.aps.org/doi/10.1103/PhysRevD.70.043008>.
- 3765 [53] M. Honda et al. "Improvement of low energy atmospheric neutrino flux
3766 calculation using the JAM nuclear interaction model". In: *Phys. Rev. D* 83 (12
3767 2011), p. 123001. URL:
3768 <https://link.aps.org/doi/10.1103/PhysRevD.83.123001>.
- 3769 [54] A. Fasso et al. "FLUKA: status and prospects for hadronic applications". In:
3770 (2001).
- 3771 [55] Y. Ashie et al. "Measurement of atmospheric neutrino oscillation parameters by
3772 Super-Kamiokande I". In: *Physical Review D* 71.11 (2005). URL:
3773 <https://doi.org/10.1103/physrevd.71.112005>.
- 3774 [56] F Reines et al. "Evidence for high-energy cosmic-ray neutrino interactions". In:
3775 *Phys. Rev. Lett.* 15.9 (Aug. 1965), pp. 429–433.
- 3776 [57] D. Casper et al. "Measurement of atmospheric neutrino composition with the
3777 IMB-3 detector". In: *Phys. Rev. Lett.* 66 (20 1991), pp. 2561–2564. URL:
3778 <https://link.aps.org/doi/10.1103/PhysRevLett.66.2561>.
- 3779 [58] K S Hirata et al. "Observation of a small atmospheric ν_μ/ν_e ratio in Kamiokande".
3780 en. In: *Phys. Lett. B* 280.1-2 (Apr. 1992), pp. 146–152.
- 3781 [59] Z. Li et al. "Measurement of the tau neutrino cross section in atmospheric
3782 neutrino oscillations with Super-Kamiokande". In: *Physical Review D* 98.5 (2018).
3783 URL: <https://doi.org/10.1103/physrevd.98.052006>.
- 3784 [60] Kamiokande Collaboration. "Atmospheric neutrino oscillation analysis with
3785 external constraints in Super-Kamiokande I-IV". In: (2017).
- 3786 [61] T2K Collaboration. "Constraint on the matter-antimatter symmetry-violating
3787 phase in neutrino oscillations". en. In: *Nature* 580.7803 (Apr. 2020), pp. 339–344.

- 3788 [62] M A Acero et al. "First measurement of neutrino oscillation parameters using
3789 neutrinos and antineutrinos by NOvA". en. In: *Phys. Rev. Lett.* 123.15 (Oct. 2019),
3790 p. 151803.
- 3791 [63] M G Aartsen et al. "Measurement of atmospheric neutrino oscillations at 6–56
3792 GeV with IceCube DeepCore". In: *Phys. Rev. Lett.* 120.7 (Feb. 2018).
- 3793 [64] P Adamson et al. "Combined analysis of $\nu\mu$ Disappearance
3794 and $\nu\mu \rightarrow \nu e$ Appearance in MINOS using accelerator and atmospheric neutrinos".
3795 In: *Phys. Rev. Lett.* 112.19 (May 2014).
- 3796 [65] M. Sajjad Athar et al. "Status and perspectives of neutrino physics". In: *Progress*
3797 *in Particle and Nuclear Physics* 124 (2022), p. 103947. URL:
3798 <https://doi.org/10.1016/Fj.pnnp.2022.103947>.
- 3799 [66] S. B. Kim, T. Lasserre, and Y. Wang. "Reactor Neutrinos". In: *Adv. High Energy*
3800 *Phys.* 2013 (2013), pp. 1–34.
- 3801 [67] M Sajjad Athar et al. "Status and perspectives of neutrino physics". In: *Prog. Part.*
3802 *Nucl. Phys.* 124 (2022), p. 103947. arXiv: 2111.07586 [hep-ph].
- 3803 [68] K Abe et al. "First gadolinium loading to Super-Kamiokande". en. In: *Nucl.*
3804 *Instrum. Methods Phys. Res. A* 1027.166248 (Mar. 2022), p. 166248.
- 3805 [69] F. P. An et al. "Observation of Electron-Antineutrino Disappearance at Daya Bay".
3806 In: *Phys. Rev. Lett.* 108 (17 2012), p. 171803. URL:
3807 <https://link.aps.org/doi/10.1103/PhysRevLett.108.171803>.
- 3808 [70] J. K. Ahn et al. "Observation of Reactor Electron Antineutrinos Disappearance in
3809 the RENO Experiment". In: *Phys. Rev. Lett.* 108 (19 2012), p. 191802. URL:
3810 <https://link.aps.org/doi/10.1103/PhysRevLett.108.191802>.
- 3811 [71] Y. Abe et al. "Indication of Reactor $\bar{\nu}_e$ Disappearance in the Double Chooz
3812 Experiment". In: *Phys. Rev. Lett.* 108 (13 2012), p. 131801. URL:
3813 <https://link.aps.org/doi/10.1103/PhysRevLett.108.131801>.
- 3814 [72] JUNO Collaboration. *TAO Conceptual Design Report: A Precision Measurement of the*
3815 *Reactor Antineutrino Spectrum with Sub-percent Energy Resolution*. 2020. arXiv:
3816 2005.08745 [physics.ins-det].
- 3817 [73] M P Decowski. "KamLAND's precision neutrino oscillation measurements". en.
3818 In: *Nucl. Phys. B*. 908 (July 2016), pp. 52–61.
- 3819 [74] A. Gando et al. "Constraints on θ_{13} from a three-flavor oscillation analysis of
3820 reactor antineutrinos at KamLAND". In: *Phys. Rev. D* 83 (5 2011), p. 052002. URL:
3821 <https://link.aps.org/doi/10.1103/PhysRevD.83.052002>.
- 3822 [75] P. Dunne. *Latest Neutrino oscillation results from T2K*. 2020.
- 3823 [76] Particle Data Group. "Review of particle physics". en. In: *Prog. Theor. Exp. Phys.*
3824 2020.8 (Aug. 2020).
- 3825 [77] R. L. Workman et al. "Review of Particle Physics". In: *PTEP* 2022 (2022),
3826 p. 083C01.
- 3827 [78] K. Abe et al. "Precise Measurement of the Neutrino Mixing Parameter θ_{23} from
3828 Muon Neutrino Disappearance in an Off-Axis Beam". In: *Phys. Rev. Lett.* 112 (18
3829 2014), p. 181801. URL:
3830 <https://link.aps.org/doi/10.1103/PhysRevLett.112.181801>.

- 3831 [79] R. Wendell. "Three Flavor Oscillation Analysis of Atmospheric Neutrinos in
3832 Super-Kamiokande". PhD thesis. University of North Carolina, 2008.
- 3833 [80] Adam M Dziewonski and Don L Anderson. "Preliminary reference Earth model".
3834 en. In: *Phys. Earth Planet. Inter.* 25.4 (June 1981), pp. 297–356.
- 3835 [81] Y Fukuda et al. "Evidence for oscillation of atmospheric neutrinos". In: *Phys. Rev.*
3836 *Lett.* 81.8 (Aug. 1998), pp. 1562–1567.
- 3837 [82] K. Abe et al. "Calibration of the Super-Kamiokande detector". In: *Nuclear*
3838 *Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers,*
3839 *Detectors and Associated Equipment* 737 (2014), pp. 253–272. URL:
3840 <https://doi.org/10.1016/j.nima.2013.11.081>.
- 3841 [83] L. Wan. "Atmospheric Neutrino_Super-K". In: (2022). URL:
3842 <https://zenodo.org/record/6694761>.
- 3843 [84] M. Jiang. "Study of the neutrino mass hierarchy with the atmospheric neutrino
3844 data collected in Super-Kamiokande IV". PhD thesis. Kyoto University, 2019.
- 3845 [85] S Fukuda et al. "The super-kamiokande detector". en. In: *Nucl. Instrum. Methods*
3846 *Phys. Res. A* 501.2-3 (Apr. 2003), pp. 418–462.
- 3847 [86] S. Fukuda et al. "The Super-Kamiokande detector". In: *Nuclear Instruments and*
3848 *Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and*
3849 *Associated Equipment* 501.2 (2003), pp. 418 –462. eprint:
3850 <http://www.sciencedirect.com/science/article/pii/S016890020300425X>.
- 3851 [87] Y Itow et al. "The JHF-Kamioka neutrino project". In: (2001).
- 3852 [88] M Jiang et al. "Atmospheric neutrino oscillation analysis with improved event
3853 reconstruction in Super-Kamiokande IV". en. In: *Prog. Theor. Exp. Phys.* 2019.5
3854 (May 2019).
- 3855 [89] H. Kume et al. "20 inch diameter photomultiplier". In: *Nuclear Instruments and*
3856 *Methods in Physics Research* 205.3 (1983), pp. 443–449. URL:
3857 <https://www.sciencedirect.com/science/article/pii/0167508783900078>.
- 3858 [90] A. Suzuki et al. "Improvement of 20 in. diameter photomultiplier tubes". In:
3859 *Nuclear Instruments and Methods in Physics Research Section A: Accelerators,*
3860 *Spectrometers, Detectors and Associated Equipment* 329.1-2 (May 1993), pp. 299–313.
3861 URL: [https://doi.org/10.1016/0168-9002\(93\)90949-i](https://doi.org/10.1016/0168-9002(93)90949-i).
- 3862 [91] Y Nakano et al. "Measurement of the radon concentration in purified water in
3863 the Super-Kamiokande IV detector". en. In: *Nucl. Instrum. Methods Phys. Res. A*
3864 977.164297 (Oct. 2020), p. 164297.
- 3865 [92] Hamamatsu. *Hamamatsu Photonics Photomultiplier Tubes Handbook*. URL:
3866 https://www.hamamatsu.com/content/dam/hamamatsu-photonics/sites/documents/99\SALES\LIBRARY/etd/PMT_handbook_v4E.pdf.
- 3868 [93] K Abe et al. "First gadolinium loading to Super-Kamiokande". en. In: *Nucl.*
3869 *Instrum. Methods Phys. Res. A* 1027.166248 (Mar. 2022), p. 166248.
- 3870 [94] J. F. Beacom and M. R. Vagins. "Antineutrino Spectroscopy with Large Water
3871 Čerenkov Detectors". In: *Phys. Rev. Lett.* 93 (17 2004), p. 171101. URL:
3872 <https://link.aps.org/doi/10.1103/PhysRevLett.93.171101>.

- 3873 [95] Ll Marti et al. "Evaluation of gadolinium's action on water Cherenkov detector
3874 systems with EGADS". en. In: *Nucl. Instrum. Methods Phys. Res. A* 959.163549
3875 (Apr. 2020), p. 163549.
- 3876 [96] Ll Marti et al. "Evaluation of gadolinium's action on water Cherenkov detector
3877 systems with EGADS". In: (2019).
- 3878 [97] M. Vagins. *Solar/DSNB Neutrino_SK-Gd*. 2022.
- 3879 [98] J. Focht. PhD thesis. Massachusetts Institute of Technology, 2004.
- 3880 [99] T. Tanimori et al. "Design and performance of semi-custom analog IC including
3881 two TACs and two current integrators for 'Super-Kamiokande'". In: *IEEE*
3882 *Transactions on Nuclear Science* 36.1 (1989), pp. 497–501.
- 3883 [100] J. Hosaka et al. "Solar neutrino measurements in Super-Kamiokande-I". In: *Phys.*
3884 *Rev. D* 73 (11 2006), p. 112001. URL:
3885 <https://link.aps.org/doi/10.1103/PhysRevD.73.112001>.
- 3886 [101] H Nishino et al. "High-speed charge-to-time converter ASIC for the
3887 Super-Kamiokande detector". en. In: *Nucl. Instrum. Methods Phys. Res. A* 610.3
3888 (Nov. 2009), pp. 710–717.
- 3889 [102] S. Yamada et al. "Commissioning of the New Electronics and Online System for
3890 the Super-Kamiokande Experiment". In: *IEEE Transactions on Nuclear Science* 57.2
3891 (2010), pp. 428–432.
- 3892 [103] S. Yamada et al. "New online system without hardware trigger for the
3893 Super-Kamiokande experiment". In: *2007 IEEE Nuclear Science Symposium
3894 Conference Record*. Honolulu, HI, USA: IEEE, Oct. 2007.
- 3895 [104] G. Carminati. "The new wide-band solar neutrino trigger for super-kamiokande".
3896 In: *Phys. Procedia* 61 (2015), pp. 666–672.
- 3897 [105] P A Čerenkov. "Visible radiation produced by electrons moving in a medium
3898 with velocities exceeding that of light". In: *Phys. Rev.* 52.4 (Aug. 1937),
3899 pp. 378–379.
- 3900 [106] I Frank and Ig Tamm. "Coherent visible radiation of fast electrons passing
3901 through matter". In: *Selected Papers*. Berlin, Heidelberg: Springer Berlin
3902 Heidelberg, 1991, pp. 29–35.
- 3903 [107] The T2K Collaboration. "Letter of Intent: Neutrino Oscillation Experiment at
3904 JHF". In: *KEK Proposal* (2001). eprint:
3905 <http://neutrino.kek.jp/jhfnu/loi/loi.v2.030528.pdf>.
- 3906 [108] Y Itow et al. "The JHF-Kamioka neutrino project". In: (June 2001). arXiv:
3907 [hep-ex/0106019](https://arxiv.org/abs/hep-ex/0106019) [hep-ex].
- 3908 [109] The K2K Collaboration and S H Ahn. "Detection of Accelerator-Produced
3909 Neutrinos at a Distance of 250 km". In: (Feb. 2001). arXiv: [hep-ex/0103001](https://arxiv.org/abs/hep-ex/0103001)
3910 [hep-ex].
- 3911 [110] The T2K Collaboration. "Tokai-to-Kamioka (T2K) Long Baseline Neutrino
3912 Oscillation Experiment Proposal". In: *KEK Proposal* (2006). eprint: {{<http://j-parc.jp/researcher/Hadron/en/pac\0606/pdf/p11-Nishikawa.pdf>}}.
- 3914 [111] C. Bronner. *Accelerator Neutrino I_Recent results from T2K*. 2022.

- 3915 [112] K. Abe et al. "Observation of Electron Neutrino Appearance in a Muon Neutrino
3916 Beam". In: *Phys. Rev. Lett.* 112 (6 2014), p. 061802. eprint:
3917 <https://link.aps.org/doi/10.1103/PhysRevLett.112.061802>.
- 3918 [113] T. Fukuda et al. "Proposal for precise measurement of neutrino-water
3919 cross-section in NINJA physics run". Proposal for J-PARC and KEK. 2017.
- 3920 [114] T. Ovsianikova et al. "New experiment WAGASCI to measure cross sections of
3921 neutrino interactions in water and hydrocarbon using J-PARC beam". In: *Physics*
3922 of Particles and Nuclei 48.6 (2017), pp. 1014–1017. eprint:
3923 <https://doi.org/10.1134/S1063779617060478>.
- 3924 [115] M. Antonova et al. "Baby MIND: a magnetized segmented neutrino detector for
3925 the WAGASCI experiment". In: *Journal of Instrumentation* 12.07 (2017), p. C07028.
3926 eprint: <http://stacks.iop.org/1748-0221/12/i=07/a=C07028>.
- 3927 [116] K. Abe et al. "First measurement of the charged current $\bar{\nu}_\mu$ double differential
3928 cross section on a water target without pions in the final state". In: *Phys. Rev. D*
3929 102 (1 2020), p. 012007. URL:
3930 <https://link.aps.org/doi/10.1103/PhysRevD.102.012007>.
- 3931 [117] K Abe et al. "Measurements of $\bar{\nu}_\mu$ and $\bar{\nu}_\mu + \nu_\mu$ charged-current cross-sections
3932 without detected pions or protons on water and hydrocarbon at a mean
3933 anti-neutrino energy of 0.86 GeV". In: *Progress of Theoretical and Experimental*
3934 *Physics* 2021.4 (Mar. 2021). URL: <https://doi.org/10.1093/ptep/ptab014>.
- 3935 [118] K. Matsuoka et al. "Design and performance of the muon monitor for the T2K
3936 neutrino oscillation experiment". In: *Nuclear Instruments and Methods in Physics*
3937 *Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*
3938 624.3 (2010), pp. 591 –600. eprint:
3939 <http://www.sciencedirect.com/science/article/pii/S016890021002098X>.
- 3940 [119] K Abe et al. "Improved constraints on neutrino mixing from the T2K experiment
3941 with 3.13×10^{21} protons on target". en. In: *Phys. Rev. D*. 103.11 (June 2021).
- 3942 [120] Tomislav Vladisavljevic. *Predicting the T2K neutrino flux and measuring oscillation*
3943 *parameters*. 1st ed. Springer theses. Cham, Switzerland: Springer Nature, Sept.
3944 2020.
- 3945 [121] D Beavis, A Carroll, and I Chiang. "Long baseline neutrino oscillation
3946 experiment at the AGS. Physics design report". In: (Apr. 1995).
- 3947 [122] P.-A. Amaudruz et al. "The T2K fine-grained detectors". In: *Nuclear Instruments*
3948 and *Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and*
3949 *Associated Equipment* 696 (Dec. 2012), pp. 1–31. URL:
3950 <https://doi.org/10.1016/j.nima.2012.08.020>.
- 3951 [123] N. Abgrall et al. "Time projection chambers for the T2K near detectors". In:
3952 *Nuclear Instruments and Methods in Physics Research Section A: Accelerators,*
3953 *Spectrometers, Detectors and Associated Equipment* 637.1 (May 2011), pp. 25–46. URL:
3954 <https://doi.org/10.1016/j.nima.2011.02.036>.
- 3955 [124] S. Assylbekov et al. "The T2K ND280 off-axis pi-zero detector". In: *Nuclear*
3956 *Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers,*
3957 *Detectors and Associated Equipment* 686 (Sept. 2012), pp. 48–63. URL:
3958 <https://doi.org/10.1016/j.nima.2012.05.028>.

- 3959 [125] D Allan et al. "The electromagnetic calorimeter for the T2K near detector
3960 ND280". In: *Journal of Instrumentation* 8.10 (2013), P10019–P10019. URL:
3961 <https://doi.org/10.1088%2F1748-0221%2F8%2F10%2Fp10019>.
- 3962 [126] *UA1 magnet sets off for a second new life*. 2022. URL: <https://cerncourier.com/a/ua1-magnet-sets-off-for-a-second-new-life/>.
- 3964 [127] F. Vannucci. "The NOMAD Experiment at CERN". In: *Advances in High Energy
3965 Physics* 2014 (2014), pp. 1–20. URL: <https://doi.org/10.1155/2014/129694>.
- 3966 [128] S. Aoki et al. "The T2K Side Muon Range Detector (SMRD)". In: *Nuclear
3967 Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers,
3968 Detectors and Associated Equipment* 698 (Jan. 2013), pp. 135–146. URL:
3969 <https://doi.org/10.1016/j.nima.2012.10.001>.
- 3970 [129] K. Suzuki et al. "Measurement of the muon beam direction and muon flux for the
3971 T2K neutrino experiment". In: *Progress of Theoretical and Experimental Physics*
3972 2015.5 (2015), pp. 53C01–0. URL: <https://doi.org/10.1093%2Fptep%2Fptv054>.
- 3973 [130] S. Brooks et al. *Handbook of Markov Chain Monte Carlo*. CRC Press, 2011.
- 3974 [131] W. R. Gilks, S. Richardson, and D. J. Spiegelhalter. *Markov Chain Monte Carlo in
3975 Practice*. Chapman & Hall/CRC Interdisciplinary Statistics, 1995.
- 3976 [132] C. Wret. "Minimising systematic uncertainties in the T2K experiment using
3977 near-detector and external data". PhD thesis. Imperial College London, 2018.
- 3978 [133] K. E. Duffy. "Measurement of the Neutrino Oscillation Parameters $\sin^2 \theta_{23}$, Δm_{32}^2 ,
3979 $\sin^2 \theta_{13}$, and δ_{CP} in Neutrino and Antineutrino Oscillation at T2K". PhD thesis.
3980 Oriel College, University of Oxford, 2016.
- 3981 [134] C. Bojeckho and A. Kabooth. "Muon neutrino disappearance simultaneous fit of
3982 ND280 and SK with Run 1+2+3 data using Markov Chain Monte Carlo Analysis".
3983 In: *T2K Technical Note* 140 (2013).
- 3984 [135] T. Bayes Rev. "An essay toward solving a problem in the doctrine of chances". In:
3985 *Phil. Trans. Roy. Soc. Lond.* 53 (1764), pp. 370–418.
- 3986 [136] A. Sztuc. "Standard and Non-Standard Neutrino-Antineutrino Oscillation
3987 Analyses and Event Reconstruction Studies using Markov Chain Monte Carlo
3988 Methods at T2K". PhD thesis. Imperial College London, 2021.
- 3989 [137] N. Metropolis et al. "Equation of State Calculations by Fast Computing
3990 Machines". In: *Journal of Chemical Physics* 21.6 (1970).
- 3991 [138] W. K. Hastings. "Monte Carlo Sampling Methods Using Markov Chains and
3992 Their Applications". In: *Biometrika* 57.1 (1970).
- 3993 [139] J. Dunkley et al. "Fast and reliable Markov chain Monte Carlo technique for
3994 cosmological parameter estimation". en. In: *Mon. Not. R. Astron. Soc.* 356.3 (Jan.
3995 2005), pp. 925–936.
- 3996 [140] Harold Jeffreys. *The Theory of Probability*. Oxford Classic Texts in the Physical
3997 Sciences. 1939.
- 3998 [141] R. E Kass and A. E Raftery. "Bayes factors". en. In: *J. Am. Stat. Assoc.* 90.430 (June
3999 1995), pp. 773–795.

- 4000 [142] T.T. Böhlen et al. "The FLUKA Code: Developments and Challenges for High
4001 Energy and Medical Applications". In: *Nuclear Data Sheets* 120 (2014), pp. 211
4002 –214. eprint:
4003 <http://www.sciencedirect.com/science/article/pii/S0090375214005018>.
- 4004 [143] R. Brun et al. *GEANT: Detector Description and Simulation Tool; Oct 1994*. CERN
4005 Program Library. Long Writeup W5013. Geneva: CERN, 1993. eprint:
4006 <http://cds.cern.ch/record/1082634>.
- 4007 [144] K. Abe et al. "T2K neutrino flux prediction". In: *Phys. Rev. D* 87 (1 2013),
4008 p. 012001. URL: <https://link.aps.org/doi/10.1103/PhysRevD.87.012001>.
- 4009 [145] C. Zeitnitz and T.A. Gabriel. "The GEANT-CALOR interface and benchmark
4010 calculations of ZEUS test calorimeters". In: *Nuclear Instruments and Methods in
4011 Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated
4012 Equipment* 349.1 (1994), pp. 106 –111. eprint:
4013 <http://www.sciencedirect.com/science/article/pii/0168900294906130>.
- 4014 [146] A. Fiorentini et al. "Flux Prediction and Uncertainty Updates with NA61 2009
4015 Thin Target Data and Negative Focussing Mode Predictions". In: *T2K Technical
4016 Note* 217 (2017).
- 4017 [147] N. Abgrall et al. "Measurements of cross sections and charged pion spectra in
4018 proton-carbon interactions at 31 GeV/ c ". In: *Physical Review C* 84.3 (2011). URL: <https://doi.org/10.1103/physrevc.84.034604>.
- 4019 xmlns:mml="http://www.w3.org/1998/Math/MathML"
4020 display="inline">><mml:mi>c</mml:mi></mml:math>".
- 4021
- 4022 [148] N. Abgrall et al. "Measurement of production properties of positively charged
4023 kaons in proton-carbon interactions at 31 GeV/ c ". In: *Physical Review C* 85.3 (2012). URL: <https://doi.org/10.1103/physrevc.85.035210>.
- 4024 xmlns:mml="http://www.w3.org/1998/Math/MathML"
4025 display="inline">><mml:mi>c</mml:mi></mml:math>".
- 4026
- 4027 [149] N. Abgrall et al. "Pion emission from the T2K replica target: Method, results and
4028 application". In: *Nuclear Instruments and Methods in Physics Research Section A:
4029 Accelerators, Spectrometers, Detectors and Associated Equipment* 701 (2013), pp. 99
4030 –114. eprint:
4031 <http://www.sciencedirect.com/science/article/pii/S016890021201234X>.
- 4032 [150] M. Apollonio et al. "Forward production of charged pions with incident protons
4033 on nuclear targets at the CERN Proton Synchrotron". In: *Phys. Rev. C* 80 (3 2009),
4034 p. 035208. eprint: <https://link.aps.org/doi/10.1103/PhysRevC.80.035208>.
- 4035 [151] B. Blau et al. "The superconducting magnet of AMS-02". In: *Nuclear Physics B -
4036 Proceedings Supplements* 113.1-3 (Dec. 2002), pp. 125–132. URL:
4037 [https://doi.org/10.1016/s0920-5632\(02\)01831-5](https://doi.org/10.1016/s0920-5632(02)01831-5).
- 4038 [152] S. Haino et al. "Measurements of primary and atmospheric cosmic-ray spectra
4039 with the BESS-TeV spectrometer". In: *Physics Letters B* 594.1-2 (July 2004),
4040 pp. 35–46. URL: <https://doi.org/10.1016/j.physletb.2004.05.019>.
- 4041 [153] NASA. *U.S. Standard Atmosphere, 1976*. 1976.

- 4042 [154] S. Roesler, R. Engel, and J. Ranft. "The Monte Carlo Event Generator
4043 DPMJET-III". In: *Advanced Monte Carlo for Radiation Physics, Particle Transport*
4044 *Simulation and Applications*. Springer Berlin Heidelberg, 2001, pp. 1033–1038. URL:
4045 <https://doi.org/10.1007/978-3-642-18211-2\166>.
- 4046 [155] Koji Niita et al. "PHITS—a particle and heavy ion transport code system". In:
4047 *Radiation Measurements* 41.9-10 (Oct. 2006), pp. 1080–1090. URL:
4048 <https://doi.org/10.1016/j.radmeas.2006.07.013>.
- 4049 [156] T Sanuki et al. "Measurements of atmospheric muon spectra at mountain
4050 altitude". In: *Physics Letters B* 541.3-4 (2002), pp. 234–242. URL:
4051 <https://doi.org/10.1016%2Fs0370-2693%2802%2902265-7>.
- 4052 [157] P. Achard et al. "Measurement of the atmospheric muon spectrum from 20 to
4053 3000 GeV". In: *Physics Letters B* 598.1-2 (2004), pp. 15–32. URL:
4054 <https://doi.org/10.1016%2Fj.physletb.2004.08.003>.
- 4055 [158] K. Sato. *Atmospheric Neutrino_Reviews on neutrino fluxes (low E atm nu)*. 2022.
- 4056 [159] Y. Hayato and L. Pickering. "The NEUT neutrino interaction simulation program
4057 library". In: *The European Physical Journal Special Topics* 230.24 (Oct. 2021),
4058 pp. 4469–4481. URL: <https://doi.org/10.1140/epjs/s11734-021-00287-7>.
- 4059 [160] Y. Hayato. "A Neutrino Interaction Simulation Program Library NEUT". In: *Acta
4060 Physica Polonica B* 40.9 (2009).
- 4061 [161] C.H. Llewellyn Smith. "Neutrino reactions at accelerator energies". In: *Physics
4062 Reports* 3.5 (1972), pp. 261 –379. eprint:
4063 <http://www.sciencedirect.com/science/article/pii/0370157372900105>.
- 4064 [162] O. Benhar, A. Fabrocini, and S. Fantoni. "The nucleon spectral function in infinite
4065 nuclear matter". In: *Nuclear Physics A* 497 (June 1989), pp. 423–430. URL:
4066 [https://doi.org/10.1016/0375-9474\(89\)90484-3](https://doi.org/10.1016/0375-9474(89)90484-3).
- 4067 [163] R. Bradford et al. "A New Parameterization of the Nucleon Elastic Form Factors".
4068 In: *Nuclear Physics B - Proceedings Supplements* 159 (2006). Proceedings of the 4th
4069 International Workshop on Neutrino-Nucleus Interactions in the Few-GeV
4070 Region, pp. 127 –132. eprint:
4071 <http://www.sciencedirect.com/science/article/pii/S0920563206005184>.
- 4072 [164] A. A. Aguilar-Arevalo et al. "First measurement of the muon neutrino charged
4073 current quasielastic double differential cross section". In: *Physical Review D* 81.9
4074 (2010). URL: <https://doi.org/10.1103%2Fphysrevd.81.092005>.
- 4075 [165] R. Gran et al. "Neutrino-nucleus quasi-elastic and 2p2h interactions up to 10
4076 GeV". In: *Phys. Rev. D* 88 (11 2013), p. 113007. eprint:
4077 <https://link.aps.org/doi/10.1103/PhysRevD.88.113007>.
- 4078 [166] Ch. Berger and L. M. Sehgal. "Lepton mass effects in single pion production by
4079 neutrinos". In: *Phys. Rev. D* 76 (11 2007), p. 113004. URL:
4080 <https://link.aps.org/doi/10.1103/PhysRevD.76.113004>.
- 4081 [167] Ch. Berger and L. M. Sehgal. "Partially conserved axial vector current and
4082 coherent pion production by low energy neutrinos". In: *Phys. Rev. D* 79 (5 2009),
4083 p. 053003. eprint: <https://link.aps.org/doi/10.1103/PhysRevD.79.053003>.

- 4084 [168] T. Sjöstrand. "High-energy-physics event generation with PYTHIA 5.7 and
4085 JETSET 7.4". In: *Computer Physics Communications* 82.1 (Aug. 1994), pp. 74–89.
4086 URL: [https://doi.org/10.1016/0010-4655\(94\)90132-5](https://doi.org/10.1016/0010-4655(94)90132-5).
- 4087 [169] C Bronner and M Hartz. "Tuning of the Charged Hadrons Multiplicities for Deep
4088 Inelastic Interactions in NEUT". In: *Proceedings of the 10th International Workshop*
4089 *on Neutrino-Nucleus Interactions in Few-GeV Region (NuInt15)*. Journal of the
4090 Physical Society of Japan, Dec. 2016. URL:
4091 <https://doi.org/10.7566/jpsc.12.010041>.
- 4092 [170] M. Glück, E. Reya, and A. Vogt. "Dynamical parton distributions revisited". In:
4093 *The European Physical Journal C* 5.3 (1998), pp. 461–470. URL:
4094 <https://doi.org/10.1007%2Fs100529800978>.
- 4095 [171] A. Bodek and U. Yang. *Axial and Vector Structure Functions for Electron- and*
4096 *Neutrino- Nucleon Scattering Cross Sections at all Q^2 using Effective Leading order*
4097 *Parton Distribution Functions*. 2010. URL: <https://arxiv.org/abs/1011.6592>.
- 4098 [172] A. Bodek and U. Yang. "Update to the Bodek-Yang Unified Model for Electron-
4099 and Neutrino- Nucleon Scattering Cross Sections". In: (2010). URL:
4100 <https://arxiv.org/abs/1012.0261>.
- 4101 [173] S. Gollapinni. "Neutrino Cross section Future". In: (2016). URL:
4102 <https://arxiv.org/abs/1602.05299>.
- 4103 [174] E. S. Pinzon Guerra et al. "Using world π^\pm -nucleus scattering data to constrain
4104 an intranuclear cascade model". In: *Phys. Rev. D* 99 (5 2019), p. 052007. URL:
4105 <https://link.aps.org/doi/10.1103/PhysRevD.99.052007>.
- 4106 [175] S. Agostinelli et al. "GEANT4: A Simulation toolkit". In: *Nucl. Instrum. Meth.*
4107 A506 (2003), pp. 250–303.
- 4108 [176] R. Brun et al. "GEANT3". In: (Sept. 1987).
- 4109 [177] A. Himmel et al. "Super-Kamiokande events and data quality studies for T2K
4110 Runs 5 and 6". In: *T2K Technical Note* 219 (2015).
- 4111 [178] S. Berkman et al. "fiTQun: A New Reconstruction Algorithm for Super-K". In:
4112 *T2K Technical Note* 146 (2013).
- 4113 [179] R.B. Patterson et al. "The extended-track event reconstruction for MiniBooNE".
4114 In: *Nuclear Instruments and Methods in Physics Research Section A: Accelerators,*
4115 *Spectrometers, Detectors and Associated Equipment* 608.1 (2009), pp. 206–224. URL:
4116 <https://doi.org/10.1016\%2Fj.nima.2009.06.064>.
- 4117 [180] M. Shiozawa. "Reconstruction algorithms in the Super-Kamiokande large water
4118 Cherenkov detector". In: *Nuclear Instruments and Methods in Physics Research*
4119 *Section A: Accelerators, Spectrometers, Detectors and Associated Equipment* 433.1-2
4120 (Aug. 1999), pp. 240–246. URL:
4121 [https://doi.org/10.1016/s0168-9002\(99\)00359-9](https://doi.org/10.1016/s0168-9002(99)00359-9).

- 4122 [181] K. Abe et al. "Search for <mml:math
4123 xmlns:mml="http://www.w3.org/1998/Math/MathML"
4124 display="inline"><mml:mi>C</mml:mi><mml:mi>P</mml:mi></mml:math>
4125 Violation in Neutrino and Antineutrino Oscillations by the T2K Experiment with
4126 <mml:math xmlns:mml="http://www.w3.org/1998/Math/MathML"
4127 display="inline"><mml:mn>2.2</mml:mn><mml:mo>x</mml:mo><mml:msup><mml:mn
4128 Protons on Target". In: *Physical Review Letters* 121.17 (Oct. 2018). URL:
4129 <https://doi.org/10.1103/physrevlett.121.171802>.
- 4130 [182] K. Abe et al. "Measurements of neutrino oscillation in appearance and
4131 disappearance channels by the T2K experiment with<mml:math
4132 xmlns:mml="http://www.w3.org/1998/Math/MathML"
4133 display="inline"><mml:mn>6.6</mml:mn><mml:mo>x</mml:mo><mml:mn>1</mml:mn
4134 on target". In: *Physical Review D* 91.7 (Apr. 2015). URL:
4135 <https://doi.org/10.1103/physrevd.91.072010>.
- 4136 [183] F and James. "MINUIT: Function Minimization and Error Analysis Reference
4137 Manual". In: (1998). CERN Program Library Long Writeups. URL:
4138 <https://cds.cern.ch/record/2296388>.
- 4139 [184] X. Li and M. Wilking. "FiTQun Event Selection Optimization". In: *T2K Technical
4140 Note* 319 (2017).
- 4141 [185] S. Tobayama. "An Analysis of the Oscillation of Atmospheric Neutrinos".
4142 PhD thesis. British Columbia U., 2016.
- 4143 [186] D. Barrow et al. "Super-Kamiokande Data Quality, MC, and Systematics in Run
4144 10". In: *T2K Technical Note* 399 (2020).
- 4145 [187] A. Maghrabi, A. Aldosari, and M. Almutairi. "Correlation analyses between solar
4146 activity parameters and cosmic ray muons between 2002 and 2012 at high cutoff
4147 rigidity". In: *Advances in Space Research* 68.7 (Oct. 2021), pp. 2941–2952. URL:
4148 <https://doi.org/10.1016/j.asr.2021.05.016>.
- 4149 [188] K. Abe et al. "Atmospheric neutrino oscillation analysis with external constraints
4150 in Super-Kamiokande I-IV". In: *Phys. Rev. D* 97 (7 2018), p. 072001. eprint:
4151 <https://link.aps.org/doi/10.1103/PhysRevD.97.072001>.
- 4152 [189] J. Beringer et al. "Review of Particle Physics". In: *Phys. Rev. D* 86 (1 2012),
4153 p. 010001. URL: <https://link.aps.org/doi/10.1103/PhysRevD.86.010001>.
- 4154 [190] S. Nakayama K. Iyogi and Y. Obayashi. "T2K data acquisition and FC event
4155 selection at Super-Kamiokande". In: *T2K Technical Note* 027 (2011).
- 4156 [191] L K Pik. "Study of the neutrino mass hierarchy with the atmospheric neutrino
4157 data observed in Super-Kamiokande". PhD thesis. Tokyo University, 2012.
- 4158 [192] R. Wendell et al. "Atmospheric neutrino oscillation analysis with subleading
4159 effects in Super-Kamiokande I, II, and III". In: *Phys. Rev. D* 81 (9 2010), p. 092004.
4160 URL: <https://link.aps.org/doi/10.1103/PhysRevD.81.092004>.
- 4161 [193] J. Hosaka et al. "Three flavor neutrino oscillation analysis of atmospheric
4162 neutrinos in Super-Kamiokande". In: *Phys. Rev. D* 74 (3 2006), p. 032002. URL:
4163 <https://link.aps.org/doi/10.1103/PhysRevD.74.032002>.

- 4164 [194] L. Munteanu et al. "Constraining the Flux and Cross Section Models with Data
4165 from ND280 using FGD1 and FGD2 for the 2020 Oscillation Analysis". In: *T2K*
4166 *Technical Note* 395 (2020).
- 4167 [195] P. Bartet et al. " ν_μ CC event selections in the ND280 tracker using Run 2+3+4
4168 data". In: *T2K Technical Note* 212 (2015).
- 4169 [196] W. Parker. "Constraining Systematic Uncertainties at T2K using Near Detector
4170 Data". PhD thesis. Royal Holloway University of London, 2020.
- 4171 [197] V. Berardi et al. " $\bar{\nu}_\mu$ event selection in the ND280 tracker using Run 5c and Run 6
4172 anti-neutrino beam data". In: *T2K Technical Note* 246 (2015).
- 4173 [198] B. Kirby J. Kim and M. Wilking. "Michel Electron Tagging in FGD1". In: *T2K*
4174 *Technical Note* 104 (2012).
- 4175 [199] J. Missert. "TN-318: Fit to Super-K Atmospheric Neutrino Data for Optimization
4176 of the fiTQun Fiducial Volume Cuts and Estimation of Detector Uncertainties".
4177 In: *T2K Technical Note* 318 (2017).
- 4178 [200] J. Chakrani et al. "NIWG model and uncertainties for 2021 oscillation analysis".
4179 In: *T2K Technical Note* 414 (2022).
- 4180 [201] Morgan Wascko. "T2K Status, Results, and Plans". *Neutrino* 2018. 2018.
- 4181 [202] T. Vladislavljevic et al. "Flux Prediction and Uncertainty with NA61/SHINE 2009
4182 Replica-Target Data (TN354 version 3.3)". In: *T2K Technical Note* 354 (2020).
- 4183 [203] G Ambrosini et al. "K/ π production ratios from 450 GeV/c protons on
4184 beryllium". en. In: *Phys. Lett. B* 420.1-2 (Feb. 1998), pp. 225–232.
- 4185 [204] E. Atkin et al. "NIWG model and uncertainties for 2019-2020 oscillation analysis".
4186 In: *T2K Technical Note* 344 (2019).
- 4187 [205] D. Barrow et al. "Flux and interaction models for the initial T2K-SK atmospheric
4188 joint fit studies". In: *T2K Technical Note* 422 (2022).
- 4189 [206] D. Barrow et al. "SK atmospheric T2K beam joint fit technical note, MaCh3
4190 details". In: *T2K Technical Note* 426 (2022).
- 4191 [207] A. A. Aguilar-Arevalo et al. "Measurement of ν_μ and $\bar{\nu}_\mu$ induced neutral current
4192 single π^0 production cross sections on mineral oil at $E_\nu \sim \mathcal{O}(1 \text{ GeV})$ ". In: *Phys.*
4193 *Rev. D* 81 (1 2010), p. 013005. eprint:
4194 <https://link.aps.org/doi/10.1103/PhysRevD.81.013005>.
- 4195 [208] P. Perio and J. Imber. "Super-K Systematic Uncertainties for RUN1-4 Joint ν_e and
4196 ν_μ Analyses". In: *T2K Technical Note* 186 (2014).
- 4197 [209] P. Perio and J. Imber. "Update of SK ν_e systematic error for 2012a oscillation
4198 analysis". In: *T2K Technical Note* 107 (2012).
- 4199 [210] C. Vilela D. Barrow. *T2K-SK Detector Matrix Uncertainties - MaCh3 Integration*.
4200 <https://git.t2k.org/t2k-sk/t2ksk-detcovmat/-/tree/feature/MaCh3Integration>. Accessed: 22-06-2022.
- 4202 [211] R. Barlow and C. Beeston. "Fitting using finite Monte Carlo samples". en. In:
4203 *Comput. Phys. Commun.* 77.2 (Oct. 1993), pp. 219–228.
- 4204 [212] J S Conway. *Incorporating nuisance parameters in likelihoods for multisource spectra*.
4205 2011.

- 4206 [213] D. Barrow et al. "Oscillation probability calculation for the T2K+SK atmospheric
4207 joint fit". In: *T2K Technical Note* 425 (2022).
- 4208 [214] R G Calland, A C Kaboth, and D Payne. "Accelerated event-by-event neutrino
4209 oscillation reweighting with matter effects on a GPU". In: 9.04 (Apr. 2014),
4210 P04016–P04016. URL: <https://doi.org/10.1088/1748-0221/9/04/p04016>.
- 4211 [215] R. Wendell. <http://www.phy.duke.edu/raw22/public/Prob3++/>.
- 4212 [216] F. Kallenborn et al. "Massively parallel computation of atmospheric neutrino
4213 oscillations on CUDA-enabled accelerators". In: *Computer Physics Communications*
4214 234 (2019), pp. 235–244. URL: <https://www.sciencedirect.com/science/article/pii/S0010465518302790>.
- 4215 [217] L. Warsame. *MaCh3 Analysis Progress*. URL: <https://indico.fnal.gov/event/50217/contributions/241232/attachments/155318/202209/MaCh3ProgressforDUNELBL\underscoreMay17%20%282%29.pdf>.
- 4216 [218] S. Bourret et al. "Neutrino oscillation tomography of the Earth with
4217 KM3NeT-ORCA". In: *Journal of Physics: Conference Series* 888 (2017), p. 012114.
4218 URL: <https://doi.org/10.1088/2F1742-6596%2F888%2F1%2F012114>.
- 4219 [219] C. Rott, A. Taketa, and D. Bose. "Spectrometry of the Earth using Neutrino
4220 Oscillations". In: *Scientific Reports* 5.1 (Oct. 2015). URL:
4221 <https://doi.org/10.1038/srep15225>.
- 4222 [220] K. Hagiwara, N. Okamura, and K. Senda. "The earth matter effects in neutrino
4223 oscillation experiments from Tokai to Kamioka and Korea". In: *Journal of High
4224 Energy Physics* 2011.9 (Sept. 2011). URL:
4225 [https://doi.org/10.1007/jhep09\(2011\)082](https://doi.org/10.1007/jhep09(2011)082).
- 4226 [221] D. Typinski. *Earth Gravity*.
4227 <http://www.typnet.net/Essays/EarthGravGraphics/EarthGrav.pdf>.
4228 Accessed: 24-06-2022.
- 4229 [222] D. Barrow. *T2K Beam + SK Atmospheric Joint Fit*. 2022.
- 4230 [223] E. Atkin et al. "Measuring PMNS parameters in a joint ND280-SK analysis using
4231 MCMC". In: *T2K Technical Note* 393 (2020).