

1      **The Sensitivity to Oscillation Parameters**  
2      **from a Simultaneous Beam and**  
3      **Atmospheric Neutrino Analysis that**  
4      **combines the T2K and SK Experiments**



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10     *Doctor of Philosophy*  
11     Michaelmas 2022

# Abstract

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# Acknowledgements

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# 1

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126

## Introduction

# 2

<sup>127</sup>

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## Neutrino Oscillation Physics

<sup>129</sup> When first proposed, neutrinos were expected to be approximately massless  
<sup>130</sup> fermions that only interact through weak and gravitational forces. This meant  
<sup>131</sup> they were very difficult to detect as they can pass through significant amounts  
<sup>132</sup> of matter without interacting. Despite this, experimental neutrino physics has  
<sup>133</sup> developed many different detection techniques and observed neutrinos from  
<sup>134</sup> both natural and artificial sources. In direct tension with Standard Model physics,  
<sup>135</sup> neutrinos have been determined to oscillate between different lepton flavours,  
<sup>136</sup> requiring them to have mass.

<sup>137</sup> The observation techniques which led to the discovery of the neutrino are doc-  
<sup>138</sup> umented in section 2.1. The theory underpinning neutrino oscillation is described  
<sup>139</sup> in section 2.2 and includes the approximations which can be made to simplify  
<sup>140</sup> the understanding of neutrino oscillation in the two-flavour approximation. Past,  
<sup>141</sup> current, and future neutrino experiments are detailed in section 2.3, including the  
<sup>142</sup> reactor, atmospheric, and long-baseline accelerator neutrino sources that have  
<sup>143</sup> been used to successfully constrain oscillation parameters. Finally, the current  
<sup>144</sup> state of oscillation parameter measurements are summarised in section 2.4.

## 2.1 Discovery of Neutrinos

At the start of the 20<sup>th</sup> century, the electrons emitted from the  $\beta$ -decay of the nucleus were found to have a continuous energy spectrum [1, 2]. This observation seemingly broke the energy conservation invoked within that period's nuclear models. In 1930, Pauli provided a solution to this problem in the form of a new particle, the neutrino (originally termed "neutron"). It was theorized to be an electrically neutral spin-1/2 fermion with a mass smaller than that of the electron [3]. This neutrino was emitted with the electron in  $\beta$ -decay to alleviate the apparent breaking of energy conservation. As a predecessor of today's weak interaction model, Fermi's theory of  $\beta$ -decay developed the understanding by coupling the four constituent particles: electron, proton, neutron, and neutrino, into a quantitative model [4].

Whilst Pauli was not convinced of the ability to detect neutrinos, the first observations of the particle were made in the mid-1950s when neutrinos from a reactor were observed via the inverse  $\beta$ -decay (IBD) process,  $\bar{\nu}_e + p \rightarrow n + e^+$  [5, 6]. The detector consisted of two parts: a neutrino interaction medium and a liquid scintillator. The interaction medium was built from two water tanks, loaded with cadmium chloride to allow for increased efficiency in the detection of neutron capture. The positron emitted from IBD annihilates,  $e^+ + e^- \rightarrow 2\gamma$ , generating a prompt signal and the neutron is captured on the cadmium via  $n + ^{108}Cd \rightarrow ^{109*}Cd \rightarrow ^{109}Cd + \gamma$ , producing a delayed signal. An increase in the coincidence rate was observed when the reactor was operating which was interpreted as interactions from neutrinos generated in the reactor.

After the discovery of the  $\nu_e$ , the question of how many flavours of neutrino exist was asked. In 1962, a measurement of the  $\nu_\mu$  was conducted at the Brookhaven National Laboratory [7]. A proton beam was directed at a beryllium target, generating pions which then decayed via  $\pi^\pm \rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu)$ , and the subsequent interactions of the  $\nu_\mu$  were observed. As the subsequent interaction of the neutrino generated muons rather than electrons, it was determined that

<sup>174</sup> the  $\nu_\mu$  was fundamentally different from  $\nu_e$ . The final observation to be made  
<sup>175</sup> was that of the  $\nu_\tau$  from the DONUT experiment [8]. Three neutrinos seem the  
<sup>176</sup> obvious solution as it mirrors the known number of charged leptons (as they form  
<sup>177</sup> weak isospin doublets) but there could be evidence of more. Several neutrino  
<sup>178</sup> experiments have found anomalous results [9, 10] which could be attributed  
<sup>179</sup> to “sterile” neutrinos. These hypothesised particles are not affected by gauge  
<sup>180</sup> interactions in the Standard Model so their presence can only be inferred through  
<sup>181</sup> the observation of non-standard oscillation modes. However, cosmological  
<sup>182</sup> observations indicate the number of neutrino species  $N_{eff} = 2.99 \pm 0.17$  [11], as  
<sup>183</sup> measured from the cosmic microwave background power spectrum. LEP also  
<sup>184</sup> measured the number of active neutrino flavours to be  $N_\nu = 2.9840 \pm 0.0082$  [12]  
<sup>185</sup> from measurements of the Z-decay width, but this does not strongly constrain  
<sup>186</sup> the number of sterile neutrinos.

## <sup>187</sup> 2.2 Theory of Neutrino Oscillation

<sup>188</sup> A neutrino generated with lepton flavour  $\alpha$  can change into a different lepton  
<sup>189</sup> flavour  $\beta$  after propagating some distance. This phenomenon is called neutrino  
<sup>190</sup> oscillation and requires that neutrinos must have a non-zero mass. This behaviour  
<sup>191</sup> has been characterised by the Pontecorvo-Maki-Nakagawa-Sakata (PMNS) [13–  
<sup>192</sup> 15] mixing matrix which describes how the flavour and mass of neutrinos are  
<sup>193</sup> associated. This is analogous to the Cabibbo-Kobayashi-Maskawa (CKM) [16]  
<sup>194</sup> matrix measured in quark physics.

### <sup>195</sup> 2.2.1 Three Flavour Oscillations

<sup>196</sup> The PMNS parameterisation defines three flavour eigenstates,  $\nu_e$ ,  $\nu_\mu$  and  $\nu_\tau$   
<sup>197</sup> (indexed  $\nu_\alpha$ ), which are eigenstates of the weak interaction and three mass  
<sup>198</sup> eigenstates,  $\nu_1$ ,  $\nu_2$  and  $\nu_3$  (indexed  $\nu_i$ ). Each mass eigenstate is the superposition

<sup>199</sup> of all three flavour states,

$$|\nu_i\rangle = \sum_{\alpha} U_{\alpha i} |\nu_{\alpha}\rangle. \quad (2.1)$$

<sup>200</sup> Where  $U$  is the  $3 \times 3$  PMNS matrix which is unitary and connects the mass  
<sup>201</sup> and flavour eigenstates.

<sup>202</sup> The weak interaction, when interacting via a  $W^{\pm}$  boson, couples to flavour  
<sup>203</sup> eigenstates so neutrinos interact with leptons of the same flavour. The prop-  
<sup>204</sup> agation of a neutrino flavour eigenstate, in a vacuum, can be re-written as a  
<sup>205</sup> plane-wave solution to the time-dependent Schrödinger equation,

$$|\nu_{\alpha}(t)\rangle = \sum_i U_{\alpha i}^{*} |\nu_i\rangle e^{-i\phi_i}. \quad (2.2)$$

<sup>206</sup> The  $\phi_i$  term can be expressed in terms of the energy,  $E_i$ , and magnitude of the  
<sup>207</sup> three momenta,  $p_i$ , of the neutrino,  $\phi_i = E_i t - p_i x$  ( $t$  and  $x$  being time and position  
<sup>208</sup> coordinates). The probability of observing a neutrino of flavour eigenstate  $\beta$  from  
<sup>209</sup> one which originated as flavour  $\alpha$  can be calculated as,

$$P(\nu_{\alpha} \rightarrow \nu_{\beta}) = |\langle \nu_{\beta} | \nu_{\alpha}(t) \rangle|^2 = \sum_{i,j} U_{\alpha i}^{*} U_{\beta i} U_{\alpha j} U_{\beta j}^{*} e^{-i(\phi_j - \phi_i)}. \quad (2.3)$$

<sup>210</sup> The term within the exponential can be represented as,

$$\phi_j - \phi_i = E_j t - E_i t - p_j x + p_i x. \quad (2.4)$$

<sup>211</sup> For a relativistic particle,  $E_i \gg m_i$ , a Taylor series expansion means,

$$p_i = \sqrt{E_i^2 - m_i^2} \approx E_i - \frac{m_i^2}{2E_i}. \quad (2.5)$$

<sup>212</sup> Making the approximations that neutrinos are relativistic, the mass eigenstates  
<sup>213</sup> were created with the same energy and that  $x = L$ , where  $L$  is the distance  
<sup>214</sup> travelled by the neutrino, Equation 2.4 then becomes

$$\phi_j - \phi_i = \frac{\Delta m_{ij}^2 L}{2E}, \quad (2.6)$$

where  $\Delta m_{ij}^2 = m_i^2 - m_j^2$ . This, combined with further use of unitarity relations results in Equation 2.3 becoming

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\beta) &= \delta_{\alpha\beta} - 4 \sum_{i>j} \Re \left( U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin^2 \left( \frac{\Delta m_{ij}^2 L}{4E} \right) \\ &\quad + (-) 2 \sum_{i>j} \Im \left( U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin \left( \frac{\Delta m_{ij}^2 L}{2E} \right). \end{aligned} \quad (2.7)$$

Where  $\delta_{\alpha\beta}$  is the Kronecker delta function and the negative sign on the last term is included for the oscillation probability of antineutrinos. As an important point to note, the observation of oscillation probability requires a non-zero value of  $\Delta m_{ij}^2$ , which in turn requires that neutrinos have differing masses.

Typically, the PMNS matrix is parameterised into three mixing angles, a charge parity (CP) violating phase  $\delta_{CP}$ , and two Majorana phases  $\alpha_{1,2}$ ,

$$U = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\text{Atmospheric, Accelerator}} \underbrace{\begin{pmatrix} c_{13} & 0 & s_{13} e^{-i\delta_{CP}} \\ 0 & 1 & 0 \\ -s_{13} e^{-i\delta_{CP}} & 0 & c_{13} \end{pmatrix}}_{\text{Reactor, Accelerator}} \times \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Reactor, Solar}} \underbrace{\begin{pmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Majorana}}. \quad (2.8)$$

Where  $s_{ij} = \sin(\theta_{ij})$  and  $c_{ij} = \cos(\theta_{ij})$ . The oscillation parameters are often grouped: (1,2) as “solar”, (2,3) as “atmospheric” and (1,3) as “reactor”. Many neutrino experiments aim to measure the PMNS parameters from a wide array of origins, as is the purpose of this thesis.

The Majorana phase,  $\alpha_{1,2}$ , included within the fourth matrix in Equation 2.8 is only included for completeness. For an oscillation analysis experiment, any terms containing this phase disappear due to taking the expectation value of the PMNS matrix. Measurements of these phases can be performed by experiments searching for neutrino-less double  $\beta$ -decay [17].

232 A two-flavour approximation can be obtained when one assumes the third  
233 mass eigenstate is degenerate with another. This results in the two-flavour  
234 approximation being reasonable for understanding the features of the oscillation.  
235 In this two-flavour case, the mixing matrix becomes,

$$U_{\text{2 Flav.}} = \begin{pmatrix} \cos(\theta) & \sin(\theta) \\ -\sin(\theta) & \cos(\theta) \end{pmatrix}. \quad (2.9)$$

236 This culminates in the oscillation probability,

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\alpha) &= 1 - \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right), \\ P(\nu_\alpha \rightarrow \nu_\beta) &= \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right). \end{aligned} \quad (2.10)$$

237 Where  $\alpha \neq \beta$ . For a fixed neutrino energy, the oscillation probability is  
238 a sinusoidal function depending upon the distance over which the neutrino  
239 propagates. The frequency and amplitude of oscillation are dependent upon  
240  $\Delta m^2/4E$  and  $\sin^2 2\theta$ , respectively. The oscillation probabilities presented thus far  
241 assume  $c = 1$ , where  $c$  is the speed of light in a vacuum. In more familiar units, the  
242 maximum oscillation probability for a fixed value of  $\theta$  is given at  $L[\text{km}] / E[\text{GeV}] \sim$   
243  $1.27/\Delta m^2$ . It is this calculation that determines the best  $L/E$  value for a given  
244 experiment to be designed around for measurements of a specific value of  $\Delta m^2$ .

### 245 2.2.2 The MSW Effect

246 The theory of neutrino oscillation in a vacuum has been described in subsec-  
247 tion 2.2.1. However, the beam neutrinos and atmospheric neutrinos originating  
248 from below the horizon propagate through the matter in the Earth. The coherent  
249 scattering of neutrinos from a material target modifies the Hamiltonian of the  
250 system which results in a change in the oscillation probability. This modification  
251 is termed the Mikheyev-Smirnov-Wolfenstein (MSW) effect [18–20]. This occurs  
252 because charged current scattering ( $\nu_e + e^- \rightarrow \nu_e + e^-$ , propagated by a W boson)  
253 only affects electron neutrinos whereas the neutral current scattering ( $\nu_l + l^- \rightarrow$

<sup>254</sup>  $\nu_l + l^-$ , propagated by a  $Z^0$  boson) interacts through all neutrino flavours equally.  
<sup>255</sup> In the two-flavour approximation, the effective mixing parameter becomes

$$\sin^2(2\theta) \rightarrow \sin^2(2\theta_m) = \frac{\sin^2(2\theta)}{(A/\Delta m^2 - \cos(2\theta))^2 + \sin^2(2\theta)}, \quad (2.11)$$

<sup>256</sup> where  $A = 2\sqrt{2}G_F N_e E$ ,  $N_e$  is the electron density of the medium and  $G_F$   
<sup>257</sup> is Fermi's constant. It is clear that there exists a value of  $A = \Delta m^2 \cos(2\theta)$  for  
<sup>258</sup>  $\Delta m^2 > 0$ , which results in a divergent mixing parameter, colloquially called the  
<sup>259</sup> matter resonance. This resonance regenerates the electron neutrino component of  
<sup>260</sup> the neutrino flux [18–20]. The density at which the resonance occurs is given by

$$N_e = \frac{\Delta m^2 \cos(2\theta)}{2\sqrt{2}G_F E}. \quad (2.12)$$

<sup>261</sup> At densities lower than this critical value, the oscillation probability will  
<sup>262</sup> be much closer to that of vacuum oscillation. For antineutrinos,  $N_e \rightarrow -N_e$   
<sup>263</sup> [21]. The resonance occurring from the MSW effect depends on the sign of  $\Delta m^2$ .  
<sup>264</sup> Therefore, any neutrino oscillation experiment which observes neutrinos and  
<sup>265</sup> antineutrinos which have propagated through matter can have some sensitivity  
<sup>266</sup> to the ordering of the neutrino mass eigenstates.

## <sup>267</sup> 2.3 Neutrino Oscillation Measurements

<sup>268</sup> As evidence of beyond Standard Model physics, the 2015 Nobel Prize in Physics  
<sup>269</sup> was awarded to the Super-Kamiokande (SK) [22] and Sudbury Neutrino Ob-  
<sup>270</sup> servatory (SNO) [23] collaborations for the first definitive observation of solar  
<sup>271</sup> and atmospheric neutrino oscillation [24]. Since then, the field has seen a wide  
<sup>272</sup> array of oscillation measurements from a variety of neutrino sources. As seen  
<sup>273</sup> in subsection 2.2.1, the neutrino oscillation probability is dependent on the ratio  
<sup>274</sup> of the propagation baseline,  $L$ , to the neutrino energy,  $E$ . It is this ratio that  
<sup>275</sup> determines the type of neutrino oscillation a particular experiment is sensitive to.

<sup>276</sup> As illustrated in Figure 2.1, there are many neutrino sources that span a  
<sup>277</sup> wide range of energies. The least energetic neutrinos are from reactor and

<sup>278</sup> terrestrial sources at  $O(1)$ MeV whereas the most energetic neutrinos originate  
<sup>279</sup> from atmospheric and galactic neutrinos of  $> O(1)$ TeV.



**Figure 2.1:** The electro-weak cross-section for  $\bar{\nu}_e + e^- \rightarrow \bar{\nu}_e + e^-$  scattering on free electrons from various natural and man-made neutrino sources, as a function of neutrino energy. Taken from [25]

### <sup>280</sup> 2.3.1 Solar Neutrinos

<sup>281</sup> Solar neutrinos are emitted from fusion reaction chains at the centre of the Sun.  
<sup>282</sup> The solar neutrino flux, given as a function of neutrino energy for different  
<sup>283</sup> fusion and decay chains, is illustrated in Figure 2.2. Whilst proton-proton fusion  
<sup>284</sup> generates the largest flux of neutrinos, the neutrinos are low energy and are  
<sup>285</sup> difficult to reconstruct due to the IBD interaction threshold of 1.8MeV [26].  
<sup>286</sup> Consequently, most experiments focus on the neutrinos from the decay of  $^8B$   
<sup>287</sup> (via  $^8B \rightarrow ^8Be^* + e^+ + \nu_e$ ), which are higher energy.

<sup>288</sup> The first measurements of solar neutrinos observed a significant reduction in  
<sup>289</sup> the event rate compared to predictions from the Standard Solar Model [28, 29]. A  
<sup>290</sup> proposed solution to this “solar neutrino problem” was  $\nu_e \leftrightarrow \nu_\mu$  oscillations in a



**Figure 2.2:** The solar neutrino flux as a function of neutrino energy for various fusion reactions and decay chains as predicted by the Standard Solar Model. Taken from [27].

291 precursory version of the PMNS model [30]. The Kamiokande [31], Gallex [32]  
 292 and Sage [33] experiments confirmed the  $\sim 0.5$  factor deficit of solar neutrinos.

293 The conclusive solution to this problem was determined by the SNO col-  
 294 laboration [34]. Using a deuterium water target to observe  ${}^8B$  neutrinos, the  
 295 event rate of charged current (CC), neutral current (NC), and elastic scattering  
 296 (ES) interactions (Given in Equation 2.13) was simultaneously measured. CC  
 297 events can only occur for electron neutrinos, whereas the NC channel is agnostic  
 298 to neutrino flavour, and the ES reaction has a small excess sensitivity for the  
 299 detection of electron neutrino interactions. This meant that there were direct  
 300 measurements of the  $\nu_e$  and  $\nu_x$  neutrino flux. It was concluded that the CC and  
 301 ES interaction rates were consistent with the deficit previously observed. Most  
 302 importantly, the NC reaction rate was only consistent with the others under the

303 hypothesis of flavour transformation.

$$\begin{aligned} \nu_e + d &\rightarrow p + p + e^- & (CC) \\ \nu_x + d &\rightarrow p + n + \nu_x & (NC) \\ \nu_x + e^- &\rightarrow \nu_x + e^- & (ES) \end{aligned} \quad (2.13)$$

304 Since the SNO measurement, many experiments have since measured the  
 305 neutrino flux of different interaction chains within the sun [35–37]. The most  
 306 recent measurement was that of CNO-cycle neutrinos which were recently  
 307 observed with  $5\sigma$  significance by the Borexino collaboration [35].

### 308 2.3.2 Accelerator Neutrinos

309 The concept of using an artificial “neutrino beam” was first realised in 1962 [38].  
 310 Since then, many experiments have adopted the same fundamental concepts.  
 311 Typically, a proton beam is aimed at a target producing charged mesons that  
 312 decay to neutrinos. The mesons can be sign-selected by the use of magnetic  
 313 focusing horns to generate a neutrino or antineutrino beam. Pions are the primary  
 314 mesons that decay and depending on the orientation of the magnetic field, a  
 315 muon (anti-)neutrino beam is generated via  $\pi^+ \rightarrow \mu^+ + \nu_\mu$  or  $\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$ .  
 316 The decay of muons and kaons results in an irreducible intrinsic electron neutrino  
 317 background. In T2K, this background contamination is  $O(< 1\%)$  [39]. There is  
 318 also an approximately  $\sim 5\%$  “wrong-sign” neutrino background of  $\bar{\nu}_\mu$  generated  
 319 via the same decays. As the beam is generated by proton interactions (rather  
 320 than anti-proton interactions), the wrong-sign component in the antineutrino  
 321 beam is larger when operating in neutrino mode.

322 Tuning the proton energy in the beam and using beam focusing techniques  
 323 allows the neutrino energy to be set to a value that maximises the disappear-  
 324 ance oscillation probability in the  $L/E$  term in Equation 2.10. This means that  
 325 accelerator experiments are typically more sensitive to the mixing parameters as  
 326 compared to a natural neutrino source. However, the disadvantage compared  
 327 to atmospheric neutrino experiments is the cost of building a facility to provide

328 high-energy neutrinos, with a high flux, which is required for longer baselines.  
 329 Consequently, there is typically less sensitivity to matter effects and the ordering  
 330 of the neutrino mass eigenstates.

331 A neutrino experiment measures

$$R(\vec{x}) = \Phi(E_\nu) \times \sigma(E_\nu) \times \epsilon(\vec{x}) \times P(\nu_\alpha \rightarrow \nu_\beta), \quad (2.14)$$

332 where  $R(\vec{x})$  is the event rate of neutrinos at position  $\vec{x}$ ,  $\Phi(E_\nu)$  is the flux of  
 333 neutrinos with energy  $E_\nu$ ,  $\sigma(E_\nu)$  is the cross-section of the neutrino interaction and  
 334  $\epsilon(\vec{x})$  is the efficiency and resolution of the detector. In order to leverage the most  
 335 out of an accelerator neutrino experiment, the flux and cross-section systematics  
 336 need to be constrained. This is typically done via the use of a “near detector”,  
 337 situated at a baseline of  $O(1)$ km. This detector observes the unoscillated neutrino  
 338 flux and constrains the parameters used within the flux and cross-section model.

339 The first accelerator experiments to precisely measure oscillation parameters  
 340 were MINOS [40] and K2K [41]. These experiments confirmed the  $\nu_\mu$  disappear-  
 341 ance seen in atmospheric neutrino experiments by finding consistent parameter  
 342 values for  $\sin^2(\theta_{23})$  and  $\Delta m_{32}^2$ . The current generation of accelerator neutrino  
 343 experiments, T2K and NO $\nu$ A extended this field by observing  $\bar{\nu}_\mu \rightarrow \bar{\nu}_e$  and lead  
 344 the sensitivity to atmospheric mixing parameters as seen in Figure 2.6 [42]. The  
 345 two experiments differ in their peak neutrino energy, baseline, and detection tech-  
 346 nique. The NO $\nu$ A experiment is situated at a baseline of 810km from the NuMI  
 347 beamline which delivers 2GeV neutrinos. The T2K neutrino beam is peaked  
 348 around 0.6GeV and propagates 295km. Additionally, the NO $\nu$ A experiment  
 349 uses functionally identical detectors (near and far) whereas T2K uses a plastic  
 350 scintillator technique at the near detector and a water Cherenkov far detector.  
 351 The future generation experiments DUNE [43] and Hyper-Kamiokande [44]  
 352 will succeed these experiments as the high-precision era of neutrino oscillation  
 353 parameter measurements develops.

354 Several anomalous results have been observed in the LSND [9] and Mini-  
 355 BooNE [10] detectors which were designed with purposefully short baselines.

356 Parts of the neutrino community attributed these results to oscillations induced  
357 by a fourth “sterile” neutrino [45] but several searches in other experiments,  
358 MicroBooNE [46] and KARMEN [47], found no hints of additional neutrino  
359 species. The solution to the anomalous results is still being determined.

### 360 2.3.3 Atmospheric Neutrinos

361 The interactions of primary cosmic ray protons in the Earth’s upper atmosphere  
362 generate showers of energetic hadrons. These are mostly pions and kaons that  
363 decay to produce a natural source of neutrinos spanning energies of MeV to  
364 TeV [48]. The main decay is via,

$$\begin{aligned} \pi^\pm &\rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu), \\ \mu^\pm &\rightarrow e^\pm + (\nu_e, \bar{\nu}_e) + (\nu_\mu, \bar{\nu}_\mu), \end{aligned} \tag{2.15}$$

365 such that for a single pion decay, three neutrinos can be produced. The  
366 atmospheric neutrino flux energy spectra as predicted by the Bartol [49], Honda  
367 [50–52], and FLUKA [53] models are illustrated in Figure 2.3. The flux distribution  
368 peaks at an energy of  $O(10)$ GeV. The uncertainties associated with these models  
369 are dominated by the hadronic production of kaon and pions as well as the  
370 primary cosmic flux.

371 Unlike long-baseline experiments which have a fixed baseline, the distance  
372 atmospheric neutrinos propagate is dependent upon the zenith angle at which  
373 they interact. This is illustrated in Figure 2.4. Neutrinos that are generated  
374 directly above the detector ( $\cos(\theta) = 1.0$ ) have a baseline equivalent to the  
375 height of the atmosphere, whereas neutrinos that interact directly below the  
376 detector ( $\cos(\theta) = -1.0$ ) have to travel a length equal to the diameter of the Earth.  
377 This means atmospheric neutrinos have a baseline that varies from  $O(20)$ km to  
378  $O(6 \times 10^3)$ km. Any neutrino generated at or below the horizon will be subject  
379 to MSW matter resonance as they propagate through the Earth.

380 Figure 2.5 highlights the neutrino flux as a function of the zenith angle for  
381 different slices of neutrino energy. For medium to high-energy neutrinos (and to



**Figure 2.3:** Left panel: The atmospheric neutrino flux for different neutrino flavours as a function of neutrino energy as predicted by the 2007 Honda model (“This work”) [50], the 2004 Honda model (“HKKM04”)[51], the Bartol model [49] and the FLUKA model [53]. Right panel: The ratio of the muon to electron neutrino flux as predicted by all the quoted models. Both figures taken from [50].



**Figure 2.4:** A diagram illustrating the definition of zenith angle as used in the Super Kamiokande experiment [54].

382 a lesser degree for low-energy neutrinos), the flux is approximately symmetric  
 383 around  $\cos(\theta) = 0$ . To the accuracy of this approximation, the systematic  
 384 uncertainties associated with atmospheric flux for comparing upward-going  
 385 and down-going neutrino cancels. This allows the down-going events, which are

386 mostly insensitive to oscillation probabilities, to act as an unoscillated prediction  
 387 (similar to a near detector in an accelerator neutrino experiment).



**Figure 2.5:** Prediction of  $\nu_e$ ,  $\bar{\nu}_e$ ,  $\nu_\mu$ ,  $\bar{\nu}_\mu$  fluxes as a function of zenith angle as calculated by the HKKM model [52]. The left, middle and right panels represent three values of neutrino energy, 0.32GeV, 1.0GeV and 3.2GeV respectively. Predictions for other models including Bartol [49], Honda [50] and FLUKA [53] are given in [54].

388 Precursory hints of atmospheric neutrinos were observed in the mid-1960s  
 389 searching for  $\nu_\mu + X \rightarrow X^* + \mu^\pm$  [55]. This was succeeded by the IMB-3 [56]  
 390 and Kamiokande [57] experiments which measured the double ratio of muon  
 391 to electron neutrinos in data to Monte Carlo,  $R(\nu_\mu/\nu_e) = (\mu/e)_{Data}/(\mu/e)_{MC}$ .  
 392 Both experiments were found to have a consistent deficit of muon neutrinos,  
 393 with  $R(\nu_\mu/\nu_e) = 0.67 \pm 0.17$  and  $R(\nu_\mu/\nu_e) = 0.658 \pm 0.016 \pm 0.035$ , respectively.  
 394 Super-Kamiokande (SK) [54] extended this analysis by fitting oscillation pa-  
 395 rameters in  $P(\nu_\mu \rightarrow \nu_\tau)$  which found best fit parameters  $\sin^2(2\theta) > 0.92$  and  
 396  $1.5 \times 10^{-3} < \Delta m^2 < 3.4 \times 10^{-3}\text{eV}^2$ .

397 Since then, atmospheric neutrino experiments have been making precision  
 398 measurements of the  $\sin^2(\theta_{23})$  and  $\Delta m^2_{32}$  oscillation parameters. Atmospheric  
 399 neutrino oscillation is dominated by  $P(\nu_\mu \rightarrow \nu_\tau)$ , where SK observed a  $4.6\sigma$   
 400 discovery of  $\nu_\tau$  appearance [58]. Figure 2.6 illustrates the current estimates on  
 401 the atmospheric mixing parameters, from a wide range of atmospheric and  
 402 accelerator neutrino observatories.



**Figure 2.6:** Constraints on the atmospheric oscillation parameters,  $\sin^2(\theta_{23})$  and  $\Delta m_{32}^2$ , from atmospheric and long-baseline experiments: SK [59], T2K [60], NOvA [61], IceCube [62] and MINOS [63]. Figure taken from [64].

### 2.3.4 Reactor Neutrinos

As illustrated in the first discovery of neutrinos (section 2.1), nuclear reactors are a very useful artificial source of electron antineutrinos. For reactors that use low-enriched uranium  $^{235}\text{U}$  as fuel, the antineutrino flux is dominated by the  $\beta$ -decay fission of  $^{235}\text{U}$ ,  $^{238}\text{U}$ ,  $^{239}\text{Pu}$  and  $^{241}\text{Pu}$  [65] as illustrated in Figure 2.7.

Due to their low energy, reactor electron antineutrinos predominantly interact via the inverse  $\beta$ -decay (IBD) interaction. The typical signature contains two signals delayed by  $O(200)\mu\text{s}$ ; firstly the prompt photons from positron annihilation, and secondly the photon emitted ( $E_{tot}^\gamma = 2.2\text{MeV}$ ) from de-excitation after neutron capture on hydrogen. Searching for both signals improves the detector's ability to distinguish between background and signal events [67].

There are many short baseline experiments ( $L \sim O(1)\text{km}$ ) that have measured the  $\sin^2(\theta_{13})$  and  $\Delta m_{32}^2$  oscillation parameters. Daya Bay [68], RENO [69] and Double Chooz [70] have all provided precise measurements, with the first discovery of a non-zero  $\theta_{13}$  made by Daya Bay and RENO (and complemented by T2K [70]). The constraints on  $\sin^2(\theta_{13})$  by the reactor experiments lead the field. They



**Figure 2.7:** Reactor electron antineutrino fluxes for  $^{235}\text{U}$  (Black),  $^{238}\text{U}$  (Green),  $^{239}\text{Pu}$  (Purple), and  $^{241}\text{Pu}$  (Orange) isotopes. The inverse  $\beta$ -decay cross-section (Blue) and corresponding measurable neutrino spectrum (Red) are also given. Top panel: Schematic of Inverse  $\beta$ -decay interaction including the eventual capture of the emitted neutron. This capture emits a  $\gamma$ -ray which provides a second signal of the event. Taken from [66].

are often used as external inputs to accelerator neutrino experiments to improve their sensitivity to  $\delta_{CP}$  and mass hierarchy determination. JUNO-TAO [71], a small collaboration within the larger JUNO experiment, is a next-generation reactor experiment that aims to precisely measure the isotopic antineutrino yields from the different fission chains.

Kamland [72] is the only experiment to have observed reactor neutrinos using a long baseline (flux weighted averaged baseline of  $L \sim 180\text{km}$ ) which allows it to have sensitivity to  $\Delta m_{21}^2$ . Combined with the SK solar neutrino experiment, the combined analysis puts the most stringent constraint on  $\Delta m_{21}^2$  [73].

## 2.4 Summary Of Oscillation Parameter Measurements

Since the first evidence of neutrino oscillations, numerous measurements of the mixing parameters have been made. Many experiments use neutrinos as a tool for the discovery of new physics (diffuse supernova background, neutrinoless double beta decay and others) so the PMNS parameters are summarised in the Particle Data Group (PDG) review tables. The analysis presented in this thesis focuses on the 2020 T2K oscillation analysis presented in [74] which the 2020 PDG constraints [75] were used. These constraints are outlined in Table 2.1.

Parameter	2020 Constraint
$\sin^2(\theta_{12})$	$0.307 \pm 0.013$
$\Delta m_{21}^2$	$(7.53 \pm 0.18) \times 10^{-5} \text{ eV}^2$
$\sin^2(\theta_{13})$	$(2.18 \pm 0.07) \times 10^{-2}$
$\sin^2(\theta_{23})$ (I.H.)	$0.547 \pm 0.021$
$\sin^2(\theta_{23})$ (N.H.)	$0.545 \pm 0.021$
$\Delta m_{32}^2$ (I.H.)	$(-2.546^{+0.034}_{-0.040}) \times 10^{-3} \text{ eV}^2$
$\Delta m_{32}^2$ (N.H.)	$(2.453 \pm 0.034) \times 10^{-3} \text{ eV}^2$

**Table 2.1:** The 2020 Particle Data Group constraints of the oscillation parameters taken from [75]. The value of  $\Delta m_{32}^2$  is given for both normal hierarchy (N.H.) and inverted hierarchy (I.H.) and  $\sin^2(\theta_{23})$  is broken down by whether its value is below (Q1) or above (Q2) 0.5.

The  $\sin^2(\theta_{13})$  measurement stems from the electron antineutrino disappearance,  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$ , and is taken as the average best-fit from the combination of Daya Bay, Reno and Double Chooz. It is often used as a prior uncertainty within other neutrino oscillation experiments, typically termed the reactor constraint. The  $\sin^2(\theta_{12})$  parameter is predominantly measured through electron neutrino disappearance,  $P(\nu_e \rightarrow \nu_{\mu,\tau})$ , in solar neutrino experiments. The long-baseline reactor neutrino experiment Kamland also has a sensitivity to this parameter and is used in a joint fit to solar data from SNO and SK, using the reactor constraint. Measurements of  $\sin^2(\theta_{23})$  are made by long-baseline and atmospheric neutrino experiments. The PDG value is a joint fit of T2K, NO $\nu$ A, MINOS and IceCube DeepCore experiments. The latest T2K-only measurement, provided at Neutrino2020 and is the basis of this thesis, is given as  $\sin^2(\theta_{23}) = 0.546^{+0.024}_{-0.046}$  [74].

The PDG constraint on  $\Delta m_{21}^2$  is provided by the KamLAND experiment using solar and geoneutrino data. This measurement utilised a  $\sin^2(\theta_{13})$  constraint from accelerator (T2K, MINOS) and reactor neutrino (Daya Bay, RENO, Double Chooz) experiments. Accelerator measurements make some of the most stringent constraints on  $\Delta m_{32}^2$  although atmospheric experiments have more sensitivity to the mass hierarchy determination. The PDG performs a joint fit of accelerator and atmospheric data, in both normal and inverted hierarchies separately. The latest T2K-only result is  $\Delta m_{32}^2 = 2.49^{+0.058}_{-0.082} \times 10^{-3} \text{ eV}^2$  favouring normal hierarchy [74]. The value of  $\delta_{CP}$  is largely undetermined. CP-conserving values of 0 and  $\pi$  were rejected with  $\sim 2\sigma$  intervals, as published in Nature, although more recent analyses have reduced the credible intervals to 90%. Since the 2020 PDG publication, there has been a new measurement of  $\sin^2(\theta_{13}) = (2.20 \pm 0.07) \times 10^{-2}$  [76], alongside updated  $\Delta m_{32}^2$  and  $\sin^2(\theta_{23})$  measurements.

Throughout this thesis, several sample spectra predictions and contours are presented, which require oscillation parameters to be assumed. Table 2.2 defines two sets of oscillation parameters, with “Asimov A” set being close to the preferred values from a previous T2K-only fit [77] and “Asimov B” being CP-conserving and further from maximal  $\theta_{23}$  mixing.

Parameter	Asimov A	Asimov B
$\Delta m_{12}^2$	$7.53 \times 10^{-5} \text{ eV}^2$	
$\Delta m_{32}^2$	$2.509 \times 10^{-3} \text{ eV}^2$	
$\sin^2(\theta_{12})$	0.304	
$\sin^2(\theta_{13})$	0.0219	
$\sin^2(\theta_{23})$	0.528	0.45
$\delta_{CP}$	-1.601	0.0

**Table 2.2:** Reference values of the neutrino oscillation parameters for two different oscillation parameter sets.

# 3

466

467

## T2K and SK Experiment Overview

468 As the successor of the Kamiokande experiment, the Super-Kamiokande (SK)  
469 collaboration has been leading atmospheric neutrino oscillation analyses for  
470 over two decades. The detector has provided some of the strongest constraints  
471 on proton decay and the first precise measurements of the  $\Delta m_{32}^2$  and  $\sin^2(\theta_{23})$   
472 neutrino oscillation parameters. The history, detection technique, and operation  
473 of the SK detector is described in section 3.1.

474 The Tokai-to-Kamioka (T2K) experiment was one of the first long-baseline  
475 experiments to use both neutrino and antineutrino beams to precisely measure  
476 charge parity violation within the neutrino sector. The T2K experiment observed  
477 the first hints of a non-zero  $\sin^2(\theta_{13})$  measurement and continues to lead the  
478 field with the constraints it provides on  $\sin^2(\theta_{13})$ ,  $\sin^2(\theta_{23})$ ,  $\Delta m_{32}^2$  and  $\delta_{CP}$ . In  
479 section 3.2, the techniques that T2K use to generate the neutrino beam and  
480 constrain systematic parameter through near detector constraints are described.

### 481 3.1 The Super-Kamiokande Experiment

482 The SK experiment began taking data in 1996 [78] and has had many modifi-  
483 cations throughout its operation. There have been seven defined periods of  
484 data taking as noted in Table 3.1. Data taking began in SK-I which ran for five

485 years. Between the SK-I and SK-II periods, approximately 55% of the PMTs were  
 486 damaged during maintenance [79]. Those that survived were equally distributed  
 487 throughout the detector in the SK-II era, which resulted in a reduced 19% photo-  
 488 coverage. From SK-III onwards, repairs to the detector meant the full suite of  
 489 PMTs was operational recovering the 40% photo-coverage. Before the start of  
 490 SK-IV, the data acquisition and electronic systems were upgraded. Between  
 491 SK-IV and SK-V, a significant effort was placed into tank open maintenance and  
 492 repair/replacement of defective PMTs in preparation for the Gadolinium upgrade;  
 493 a task for which the author of this thesis was required. Consequently, the detector  
 494 conditions were significantly changed from this point. SK-VI marked the start of  
 495 the SK-Gd era, with the detector being doped with gadolinium at a concentration  
 496 of 0.01% by concentration. SK-VII, which started during the writing of this thesis,  
 497 has increased the gadolinium concentration to 0.03% for continued operation [80].

498 The oscillation analysis presented within this thesis focuses on the SK-IV  
 499 period of running and the data taken within it. This follows from the recent  
 500 SK analysis presented in [81]. Therefore, the information presented within this  
 501 section focuses on that period.

Period	Start Date	End Date	Live-time (days)
I	April 1996	July 2001	1489.19
II	October 2002	October 2005	798.59
III	July 2006	September 2008	518.08
IV	September 2008	May 2018	3244.4
V	January 2019	July 2020	461.02
VI	July 2020	May 2022	583.3
VII	May 2022	Ongoing	N/A

502 **Table 3.1:** The various SK periods and their respective live-time. The SK-VI live-time is  
 503 calculated until 1<sup>st</sup> April 2022. SK-VII started during the writing of this thesis.

### 502 3.1.1 The SK Detector

503 The basic structure of the Super-Kamiokande (SK) detector is a cylindrical tank  
 504 with a diameter 39.3m and height 41.1m filled with ultrapure water [79]. A  
 505 diagram of the significant components of the SK detector is given in Figure 3.1.

506 The SK detector is situated in the Kamioka mine in Gifu, Japan. The mine  
 507 is underground with roughly 1km rock overburden (2.7km water equivalent  
 508 overburden) [82]. At this depth, the rate of cosmic ray muons is significantly  
 509 decreased to a value of  $\sim 2\text{Hz}$  (net rate) [83]. The top of the tank is covered  
 510 with stainless steel which is designed as a working platform for maintenance,  
 511 calibration, and location for high voltage and data acquisition electronics.



**Figure 3.1:** A schematic diagram of the Super-Kamiokande Detector. Taken from [84].

512 A smaller cylindrical structure (36.2m diameter, 33.8m height) is situated  
 513 inside the tank, with an approximate 2m gap between this structure and the outer  
 514 tank wall. The purpose of this structure is to support the photomultiplier tubes  
 515 (PMTs). The volume inside and outside the support structure is referred to as the  
 516 inner detector (ID) and outer detector (OD), respectively. In the SK-IV era, the  
 517 ID and OD are instrumented by 11,129 50cm and 1,885 20cm PMTs respectively  
 518 [79]. The ID contains a 32kton mass of water. Many analyses performed at SK  
 519 use a “fiducial volume” defined by the volume of water inside the ID excluding  
 520 some distance to the ID wall. This reduces the volume of the detector which is  
 521 sensitive to neutrino events but reduces radioactive backgrounds and allows for

522 better reconstruction performance. The nominal fiducial volume is defined as the  
523 area contained inside 2m from the ID wall for a total of 22.5kton water [85].

524 The two regions of the detector (ID and OD) are optically separated with  
525 opaque black plastic hung from the support structure. The purpose of this is  
526 to determine whether an event entered or exited the ID. This allows cosmic ray  
527 muons and partially contained events to be tagged and separated from neutrino  
528 events entirely contained within the ID. This black plastic is also used to cover  
529 the area between the ID PMTs to reduce photon reflection from the ID walls.  
530 Opposite to this, the OD is lined with a reflective material to allow photons to  
531 reflect around inside the OD until collected by one of the PMTs. Furthermore,  
532 each OD PMT is optically coupled with  $50 \times 50\text{cm}$  plates of wavelength shifting  
533 acrylic which increases the efficiency of light collection [82].

534 In the SK-IV data-taking period, the photocathode coverage of the detector, or  
535 the fraction of the ID wall instrumented with PMTs, is  $\sim 40\%$  [82]. The PMTs have  
536 a quantum efficiency (the ratio of detected electrons to incident photons) of  $\sim 21\%$   
537 for photons with wavelengths of  $360\text{nm} < \lambda < 390\text{nm}$  [86, 87]. The proportion  
538 of photoelectrons that produce a signal in the dynode of a PMT, termed the  
539 collection efficiency, is  $> 70\%$  [82]. The PMTs used within SK are most sensitive  
540 to photons with wavelength  $300\text{nm} \leq \lambda \leq 600\text{nm}$  [82]. One disadvantage of  
541 using PMTs as the detection media is that the Earth's geomagnetic field can  
542 modify its response. Therefore, a set of compensation coils is built around the  
543 inner surface of the detector to mitigate this effect [83].

544 The SK detector is filled with ultrapure water, which in a perfect world, con-  
545 tains no impurities. However, bacteria and organic compounds can significantly  
546 degrade the water quality. This decreases the attenuation length, which reduces  
547 the total number of photons that hit a PMT. To combat this, a sophisticated water  
548 treatment system has been developed [82, 88]. UV lights, mechanical filters, and  
549 membrane degasifiers are used to reduce the bacteria, suspended particulates,  
550 and radioactive materials from the water. The flow of water within the tank  
551 is also critical as it can remove stagnant bacterial growth or build-up of dust

552 on the surfaces within the tank. Gravity drifts impurities in the water towards  
553 the bottom of the tank which, if left uncontrolled, can create asymmetric water  
554 conditions between the top and bottom of the tank. Typically, the water entering  
555 the tank is cooled below the ambient temperature of the tank to control convection  
556 and inhibit bacteria growth. Furthermore, the rate of dark noise hits within PMTs  
557 is sensitive to the PMT temperature [89]. Therefore controlling the temperature  
558 gradients within the tank is beneficial for stable measurements.

559 SK-VI is the first phase of the SK experiment to use gadolinium dopants  
560 within the ultrapure water [80]. As such, the SK water system had to be replaced  
561 to avoid removing the gadolinium concentrate from the ultrapure water [90]. For  
562 an inverse  $\beta$ -decay (IBD) interaction on a water target, the emitted neutron is  
563 thermally captured on hydrogen. This process releases a 2.2MeV  $\gamma$  ray which is  
564 difficult to detect as the resulting Compton scattered electrons are very close to the  
565 Cherenkov threshold, limiting detection capability. Thermal capture of neutrons  
566 on gadolinium generates  $\gamma$  rays with higher energy (8MeV [67]) meaning they  
567 are more easily detected and reconstructed. SK-VI has 0.01% Gd loading (0.02%  
568 gadolinium sulphate by mass) which causes  $\approx$  50% of neutrons emitted by IBD to  
569 be captured on gadolinium[91, 92] . Whilst predominantly useful for low energy  
570 analyses, Gd loading allows better  $\nu/\bar{\nu}$  separation for atmospheric neutrino  
571 event selections [93]. Efforts are currently in place to increase the gadolinium  
572 concentrate to 0.03% for  $\approx$  75% neutron capture efficiency on gadolinium [94].  
573 The final stage of loading targets 0.1% concentrate for  $\approx$  90% neutron capture  
574 efficiency on gadolinium.

### 575 3.1.2 Calibration

576 The calibration of the SK detector is documented in [79] and summarised below.  
577 The analysis presented within this thesis is dependent upon ‘high energy events’  
578 (Charged particles with  $O(> 100)$ MeV momenta). These are events that are  
579 expected to generate a larger number of photons such that each PMT will  
580 be hit with multiple photons. The reconstruction of these events depends

upon the charge deposited within each PMT and the timing response of each individual PMT. Therefore, the most relevant calibration techniques to this thesis are outlined.

Before installation, 420 PMTs were calibrated to have identical charge responses and then distributed throughout the tank in a cross-shape pattern (As illustrated by Figure 3.2). These are used as a standardised measure for the rest of the PMTs installed at similar geometric positions within SK to be calibrated against. To perform this calibration, a xenon lamp is located at the center of the SK tank which flashes uniform light at 1Hz. This allows for geometrical effects, water quality variation, and timing effects to be measured in situ throughout normal data-taking periods.



**Figure 3.2:** The location of “standard PMTs” (red) inside the SK detector. Taken from [79].

When specifically performing calibration of the detector (in out-of-data taking mode), the water in the tank was circulated to avoid top/bottom asymmetric water quality. Any non-uniformity within the tank significantly affects the PMT hit probability through scattering or absorption. This becomes a dominant effect for very low-intensity light sources that are designed such that only one photon is incident upon a given PMT.

598      The gain of a PMT is defined as the ratio of the total charge of the signal  
 599      produced compared to the charge of photoelectrons emitted by the photocathodes  
 600      within the PMT. To calibrate the signal of each PMT, the “relative” and “absolute”  
 601      gain values are measured. The relative gain is the variation of gain among each  
 602      of the PMTs whereas the absolute gain is the average gain of all PMTs.

603      The relative gain is calibrated as follows. A laser is used to generate two  
 604      measurements: a high-intensity flash that illuminates every PMT with a sufficient  
 605      number of photons, and a low-intensity flash in which only a small number  
 606      of PMTs collect light. The first measurement creates an average charge,  $Q_{obs}(i)$   
 607      on PMT  $i$ , whereas the second measurement ensures that each hit PMT only  
 608      generates a single photoelectron. For the low-intensity measurement, the number  
 609      of times each PMT records a charge larger than 1/4 photoelectrons,  $N_{obs}(i)$ , is  
 610      counted. The values measured can be expressed as

$$\begin{aligned} Q_{obs}(i) &\propto I_H \times f(i) \times \epsilon(i) \times G(i), \\ N_{obs}(i) &\propto I_L \times f(i) \times \epsilon(i). \end{aligned} \tag{3.1}$$

611      Where  $I_H$  and  $I_L$  is the intensity of the high and low flashes,  $f(i)$  is the  
 612      acceptance efficiency of the  $i^{\text{th}}$  PMT,  $\epsilon(i)$  is the product of the quantum and  
 613      collection efficiency of the  $i^{\text{th}}$  PMT and  $G(i)$  is the gain of the  $i^{\text{th}}$  PMT. The relative  
 614      gain for each PMT can be determined by taking the ratio of these quantities.

615      The absolute gain calibration is performed by observing fixed energy  $\gamma$ -rays  
 616      of  $E_\gamma \sim 9\text{MeV}$  emitted isotropically from neutron capture on a NiCf source  
 617      situated at the center of the detector. This generates a photon yield of about 0.004  
 618      photoelectrons/PMT/event, meaning that  $> 99\%$  of PMT signals are generated  
 619      from single photoelectrons. A charge distribution is generated by performing  
 620      this calibration over all PMTs, and the average value of this distribution is taken  
 621      to be the absolute gain value.

622      As mentioned in subsection 3.1.1, the average quantum and collection effi-  
 623      ciency for the SK detector PMTs is  $\sim 21\%$  and  $> 70\%$  respectively. However,  
 624      these values do differ between each PMT and need to be calibrated accordingly.

625 Consequently, the NiCf source is also used to calibrate the “quantum  $\times$  collection”  
626 efficiency (denoted “QE”) value of each PMT. The NiCf low-intensity source is  
627 used as the PMT hit probability is proportional to the QE ( $N_{obs}(i) \propto \epsilon(i)$  in  
628 Equation 3.1). A Monte Carlo prediction which includes photon absorption,  
629 scattering, and reflection is made to estimate the number of photons incident on  
630 each PMT and the ratio of the number of predicted to observed hits is calculated.  
631 The difference is attributed to the QE efficiency of that PMT. This technique is  
632 extended to calculate the relative QE efficiency by normalizing the average of  
633 all PMTs which removes the dependence on the light intensity.

634 Due to differing cable lengths and readout electronics, the timing response  
635 between a photon hitting the PMT and the signal being captured by the data  
636 acquisition can be different between each PMT. Due to threshold triggers (De-  
637 scribed in subsection 3.1.3), the time at which a pulse reaches a threshold is  
638 dependent upon the size of the pulse. This is known as the ‘time-walk’ effect  
639 and also needs to be accounted for in each PMT. To calibrate the timing response,  
640 a pulse of light with width 0.2ns is emitted into the detector through a diffuser.  
641 Two-dimensional distributions of time and pulse height (or charge) are made  
642 for each PMT and are used to calibrate the timing response. This is performed  
643 in-situ during data taking with the light source pulsing at 0.03Hz.

644 The top/bottom water quality asymmetry is measured using the NiCf calibra-  
645 tion data and cross-referencing these results to the “standard PMTs”. The water  
646 attenuation length is continuously measured by the rate of vertically-downgoing  
647 cosmic-ray muons which enter via the top of the tank.

648 Dark noise is where a PMT registers a pulse that is consistent with a single  
649 photoelectron emitted from photon detection despite the PMT being in complete  
650 darkness. This is predominately caused by two processes. Firstly there is  
651 intrinsic dark noise which is where photoelectrons gain enough thermal energy  
652 to be emitted from the photocathode, and secondly, the radioactive decay of  
653 contaminants inside the structure of the PMT. Typical dark noise rate for PMTs  
654 used within SK are  $O(3)\text{kHz}$  [82]. This is lower than the expected number of

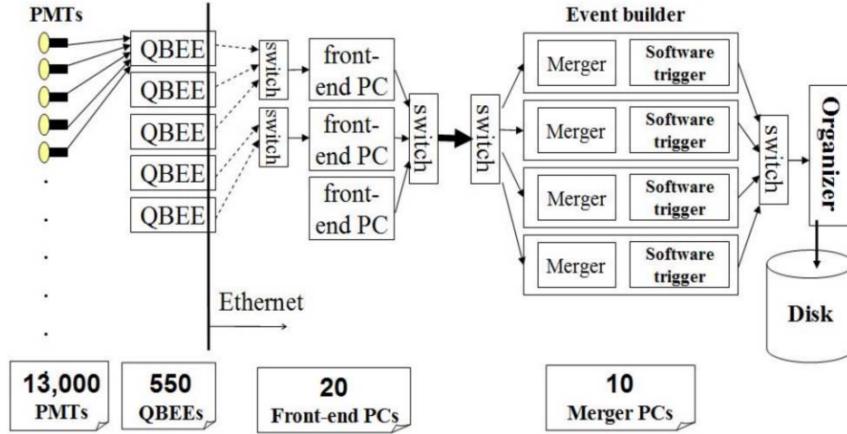
655 photons generated for a ‘high energy event’ (As described in subsection 3.1.4)  
656 but instability in this value can cause biases in reconstruction. Dark noise is  
657 related to the gain of a PMT and is calibrated using hits inside a time window  
658 recorded before an event trigger [95].

### 659 3.1.3 Data Acquisition and Triggering

660 As the analysis presented in this thesis will only use the SK-IV period of the  
661 SK experiment so this subsection focuses on the relevant points of the data  
662 acquisition and triggering systems to that SK period. The earlier data acquisition  
663 and triggering systems are documented in [96, 97].

664 Before the SK-IV period started, the existing front-end electronics were re-  
665 placed with “QTC-Based Electronics with Ethernet, QBEE” systems [98]. When  
666 the QBEE observes a signal above a 1/4 photoelectron threshold, the charge-to-  
667 time (QTC) converter generates a rectangular pulse. The start of the rectangular  
668 pulse indicates the time at which the analog photoelectron signal was received  
669 and the width of the pulse indicates the total charge integrated throughout the  
670 signal. This is then digitized by time-to-digital converters and sent to the “front-  
671 end” PCs. The digitized signal from every QBEE is then chronologically ordered  
672 and sent to the “merger” PCs. It is the merger PCs that apply the software trigger.  
673 Any triggered events are passed to the “organizer” PC. This sorts the data stream  
674 of multiple merger PCs into chronologically ordered events, which are then saved  
675 to disk. The schematic of data flow from PMTs to disk is illustrated in Figure 3.3.

676 The software trigger (described in [100]) operates by determining the number  
677 of PMT hits within a 200ns sliding window,  $N_{200}$ . This window coincides with the  
678 maximum time that a Cherenkov photon would take to traverse the length of the  
679 SK tank [97]. For lower energy events that generate fewer photons, this technique  
680 is useful for eliminating background processes like dark noise and radioactive  
681 decay which would be expected to be separated in time. When the value of  
682  $N_{200}$  exceeds some pre-defined threshold, a software trigger is issued. There are  
683 several trigger thresholds used within the SK-IV period which are detailed in



**Figure 3.3:** Schematic view of the data flow through the data acquisition and online system. Taken from [99].

684 Table 3.2. If one of these thresholds is met, the PMT hits within an extended time  
 685 window are also read out and saved to disk. In the special case of an event that  
 686 exceeds the SHE trigger but does not exceed the OD trigger, the AFT trigger looks  
 687 for delayed coincidences of 2.2MeV gamma rays emitted from neutron capture in  
 688 a  $535\mu\text{s}$  window after the SHE trigger. A similar but more complex “Wideband  
 689 Intelligent Trigger (WIT)” has been deployed and is described in [101].

Trigger	Acronym	Condition	Extended time window ( $\mu\text{s}$ )
Super Low Energy	SLE	>34/31 hits	1.3
Low Energy	LE	>47 hits	40
High Energy	HE	>50 hits	40
Super High Energy	SHE	>70/58 hits	40
Outer Detector	OD	>22 hits in OD	N/A

**Table 3.2:** The trigger thresholds and extended time windows saved around an event which were utilised throughout the SK-IV period. The exact thresholds can change and the values listed here represent the thresholds at the start and end of the SK-IV period.

### 690 3.1.4 Cherenkov Radiation

691 Cherenkov light is emitted from any highly energetic charged particle traveling  
 692 with relativistic velocity,  $\beta$ , greater than the local speed of light in a medium [102].

693 Cherenkov light is formed at the surface of a cone with a characteristic pitch angle,

$$\cos(\theta) = \frac{1}{\beta n}. \quad (3.2)$$

694 Where  $n$  is the refractive index of the medium. Consequently, the Cherenkov  
 695 momentum threshold,  $P_{thres}$ , is dependent upon the mass,  $m$ , of the charged  
 696 particle moving through the medium,

$$P_{thres} = \frac{m}{\sqrt{n^2 - 1}}. \quad (3.3)$$

697 For water, where  $n = 1.33$ , the Cherenkov threshold momentum and energy  
 698 for various particles are given in Table 3.3. In contrast,  $\gamma$ -rays are detected  
 699 indirectly via the combination of photons generated by Compton scattering  
 700 and pair production. The threshold for detection in the SK detector is typically  
 701 higher than the threshold for photon production. This is due to the fact that the  
 702 attenuation of photons in the water means that typically  $\sim 75\%$  of Cherenkov  
 703 photons reach the ID PMTs. Then the collection and quantum efficiencies  
 704 described in subsection 3.1.1 result in the number of detected photons being  
 705 lower than the number of photons which reach the PMTs.

Particle	Threshold Momentum (MeV)	Threshold Energy (MeV)
Electron	0.5828	0.7751
Muon	120.5	160.3
Pion	159.2	211.7
Proton	1070.0	1423.1

**Table 3.3:** The threshold momentum and total energy for a particle to generate Cherenkov light in ultrapure water, as calculated in Equation 3.2 in ultrapure water which has refractive index  $n = 1.33$ .

706 The Frank-Tamm equation [103] describes the relationship between the num-  
 707 ber of Cherenkov photons generated per unit length,  $dN/dx$ , the wavelength of  
 708 the photons generated,  $\lambda$ , and the relativistic velocity of the charged particle,

$$\frac{d^2N}{dxd\lambda} = 2\pi\alpha \left(1 - \frac{1}{n^2\beta^2}\right) \frac{1}{\lambda^2}. \quad (3.4)$$

709 where  $\alpha$  is the fine structure constant. For a 100MeV momentum electron,  
 710 approximately 330 photons will be produced per centimeter in the  $300\text{nm} \leq \lambda \leq$   
 711  $700\text{nm}$  region which the ID PMTs are most sensitive to [82].

## 712 3.2 The Tokai to Kamioka Experiment

713 The Tokai to Kamioka (T2K) experiment is a long-baseline neutrino oscillation  
 714 experiment located in Japan. Proposed in the early 2000s [104, 105] to replace  
 715 K2K [106], T2K was designed to observe electron neutrino appearance whilst  
 716 precisely measuring the oscillation parameters associated with muon neutrino  
 717 disappearance [107]. The experiment consists of a neutrino beam generated  
 718 at the Japan Proton Accelerator Research Complex (J-PARC), a suite of near  
 719 detectors situated 280m from the beam target, and the Super Kamiokande far  
 720 detector positioned at a 295km baseline. The cross-section view of the T2K  
 721 experiment is drawn in Figure 3.4.

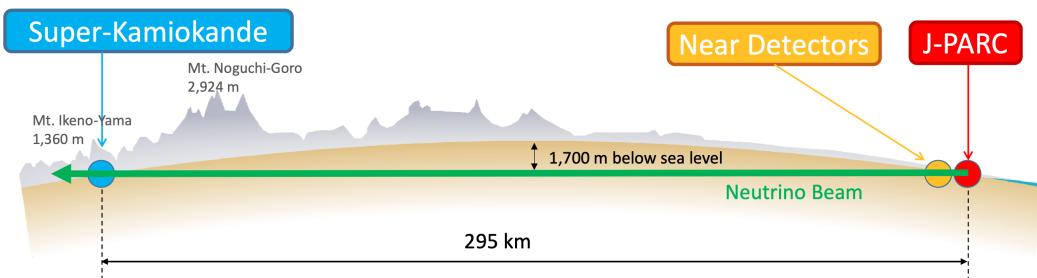
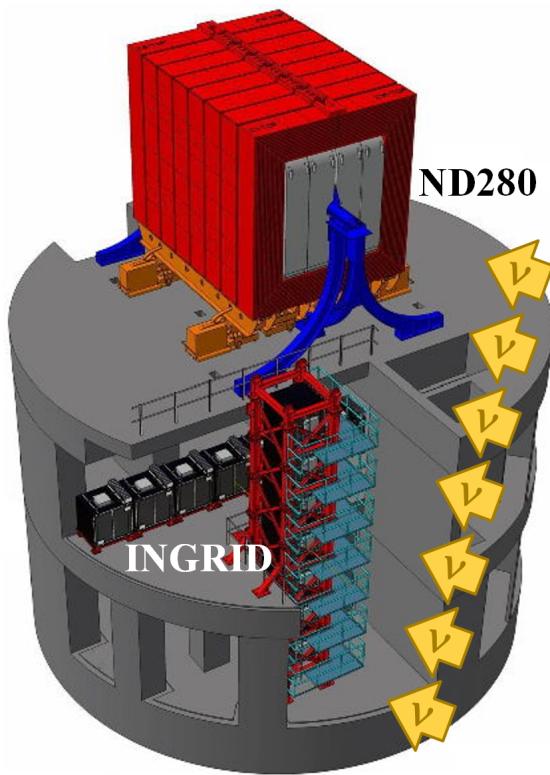


Figure 3.4: The cross-section view of the Tokai to Kamioka experiment illustrating the beam generation facility at J-PARC, the near detector situated at a baseline of 280m and the Super Kamiokande far detector situated 295km from the beam target.

722 The T2K collaboration makes world-leading measurements of the  $\sin^2(\theta_{23})$ ,  
 723  $\Delta m_{32}^2$ , and  $\delta_{CP}$  oscillation parameters. Improvements in the precision and accu-  
 724 racy of parameter estimates are still being made by including new data samples  
 725 and developing the models which describe the neutrino interactions and detector  
 726 responses [108]. Electron neutrino appearance was first observed at T2K in 2014  
 727 [109] with  $7.3\sigma$  significance.

728 The near detectors provide constraints on the beam flux and cross-section  
 729 model parameters used within the oscillation analysis by observing the unoscil-  
 730 lated neutrino beam. There are a host of detectors situated in the near detector hall  
 731 (As illustrated in Figure 3.5): ND280 (subsection 3.2.3), INGRID (subsection 3.2.4),  
 732 NINJA [110], WAGASCI [111], and Baby-MIND [112]. The latter three are not  
 733 currently used within the oscillation analysis presented in this thesis.



**Figure 3.5:** The near detector suite for the T2K experiment showing the ND280 and INGRID detectors. The distance between the detectors and the beam target is 280m.

734 Whilst this thesis presents the ND280 in terms of its purpose for the oscillation  
 735 analysis, the detector can also make many cross-section measurements at neutrino  
 736 energies of  $O(1)$ GeV for the different targets within the detector [113, 114]. These  
 737 measurements are of equal importance as they can lead the way in determining  
 738 the model parameters used in the interaction models for the future high-precision  
 739 era of neutrino physics.

### 3.2.1 Analysis Overview

There are two independent fitters, MaCh3 and BANFF, which perform the near detector fit. MaCh3 uses a bayesian Markov Chain Monte Carlo fitting technique, whereas BANFF uses a frequentist gradient descent technique. The output of each fitter is compared as a method of cross-checking the behaviour of the two fitters. This is done by comparing: the Monte Carlo predictions using various tunes, the likelihood that is calculated in each fitter and the post-fit constraint associated with every parameter used in the fit. Once validated, the output converted into a covariance matrix to describe the error and correlations between all the flux and cross-section parameters. This is then propagated to the far-detector oscillation analysis group.

The far detector group has three independent fitters: P-Theta, VALOR and MaCh3. The first two fitters use a hybrid frequentist fitting technique where the likelihood is minimised with respect to the parameters of interest and marginalised over all other parameters. These fitters use the covariance provided by the near detector fitters as a basis for implementing the near detector constraints. The MaCh3 fitter uses a simultaneous fit of all near and far detector samples. This removes any Gaussian assumptions when making the covariance matrix from the near detector results. The results for all three fitters are compares using a technique similar to the validation of the near detector fitters.

There are three particular tunes of the T2K flux and low energy cross section model typically considered. Firstly, the “generated” tune which is the set of dial values with which the Monte Carlo was generated. Secondly, the set of dial values which are taken from external data measurements and used as inputs. These are the “pre-fit” dial values. The reason these two sets of dial values are different is that the external data measurements are continually updated but it is very computationally intensive to regenerate a Monte Carlo prediction after each update. The final tune is the “post-fit”, “post-ND fit” or “post-BANFF” dial values. These are the values taken from the constraints provided by the near detector.

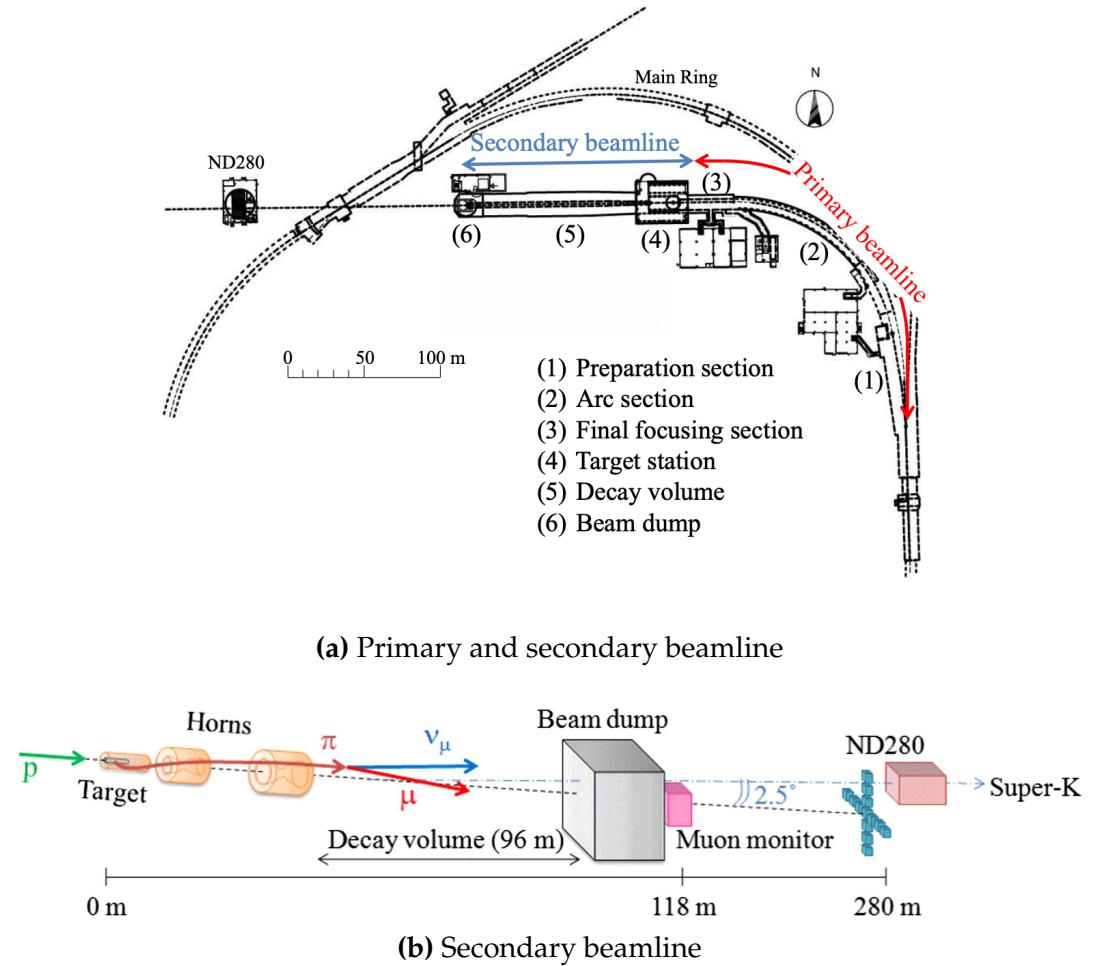
### 3.2.2 The Neutrino Beam

The neutrino beam used within the T2K experiment is described in [39, 115] and summarised below. The accelerator facility at J-PARC is composed of two sections; the primary and secondary beamlines. Figure 3.6 illustrates a schematic of the beamline, focusing mostly on the components of the secondary beamline. The primary beamline has three accelerators that progressively accelerate protons; a linear accelerator, a rapid-cycling synchrotron, and the main-ring (MR) synchrotron. Once fully accelerated by the MR, the protons have a kinetic energy of 30GeV. Eight bunches of these protons, separated by 500ns, are extracted per “spill” from the MR and directed towards a graphite target (a rod of length 91.4cm and diameter 2.6cm). Spills are extracted at 0.5Hz with  $\sim 3 \times 10^{14}$  protons contained per spill.

The secondary beamline consists of three main components: the target station, the decay volume, and the beam dump. The target station is comprised of the target, beam monitors, and three magnetic focusing horns. The proton beam interacts with the graphite target to form a secondary beam of mostly pions and kaons. The secondary beam travels through a 96m long decay volume, generating neutrinos through the following decays [39],

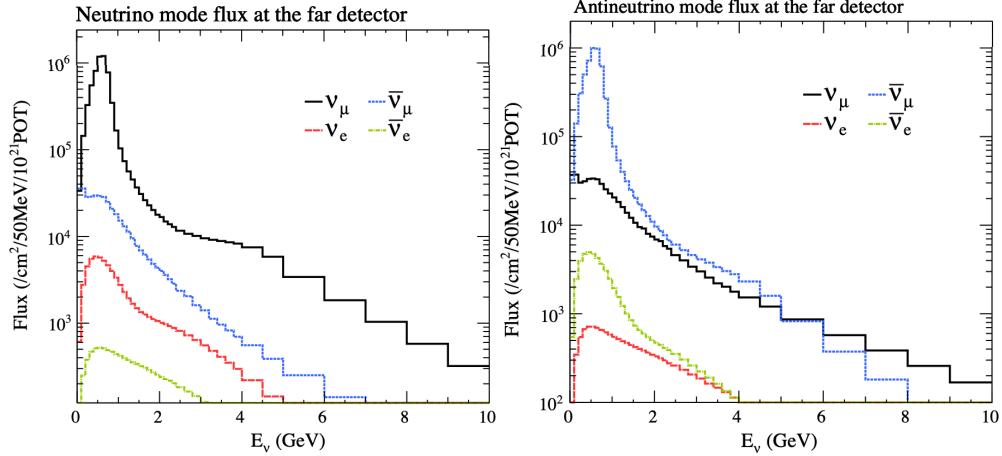
$$\begin{array}{ll}
\pi^+ \rightarrow \mu^+ + \nu_\mu & \pi^- \rightarrow \mu^- + \bar{\nu}_\mu \\
K^+ \rightarrow \mu^+ + \nu_\mu & K^- \rightarrow \mu^- + \bar{\nu}_\mu \\
\rightarrow \pi^0 + e^+ + \nu_e & \rightarrow \pi^0 + e^- + \bar{\nu}_e \\
\rightarrow \pi^0 + \mu^+ + \nu_\mu & \rightarrow \pi^0 + \mu^- + \bar{\nu}_\mu \\
K_L^0 \rightarrow \pi^- + e^+ + \nu_e & K_L^0 \rightarrow \pi^+ + e^- + \bar{\nu}_e \\
\rightarrow \pi^- + \mu^+ + \nu_\mu & \rightarrow \pi^+ + \mu^- + \bar{\nu}_\mu \\
\mu^+ \rightarrow e^+ + \bar{\nu}_\mu + \nu_e & \mu^- \rightarrow e^- + \nu_\mu + \bar{\nu}_e.
\end{array}$$

The electrically charged component of the secondary beam is focused towards the far detector by the three magnetic horns. These horns direct charged particles of a particular polarity towards SK whilst defocusing the oppositely charged particles. This allows a mostly neutrino or mostly antineutrino beam to be used within the experiment, denoted as “forward horn current (FHC)” or “reverse horn current (RHC)” respectively.



**Figure 3.6:** Top panel: Bird's eye view of the most relevant part of primary and secondary beamline used within the T2K experiment. The primary beamline is the main-ring proton synchrotron, kicker magnet, and graphite target. The secondary beamline consists of the three focusing horns, decay volume, and beam dump. Figure taken from [115]. Bottom panel: The side-view of the secondary beamline including the focusing horns, beam dump and neutrino detectors. Figure taken from [116].

795     Figure 3.7 illustrates the different contributions to the FHC and RHC neutrino  
 796     flux. The low energy flux is dominated by the decay of pions whereas kaon  
 797     decay becomes the dominant source of neutrinos for  $E_\nu > 3\text{GeV}$ . The “wrong-  
 798     sign” component, which is the  $\bar{\nu}_\mu$  background in a  $\nu_\mu$  beam, and the intrinsic  
 799     irreducible  $\nu_e$  background, are predominantly due to muon decay for  $E_\nu <$   
 800      $2\text{GeV}$ . As the antineutrino production cross-section is smaller than the neutrino  
 801     cross-section, the wrong-sign component is more dominant in the RHC beam  
 802     as compared to that in the FHC beam.



**Figure 3.7:** The Monte Carlo prediction of the energy spectrum for each flavour of neutrino ( $\nu_e$ ,  $\bar{\nu}_e$ ,  $\nu_\mu$  and  $\bar{\nu}_\mu$ ) in the neutrino dominated beam FHC mode (Left) and antineutrino dominated beam RHC mode (Right) expected at SK. Taken from [117].

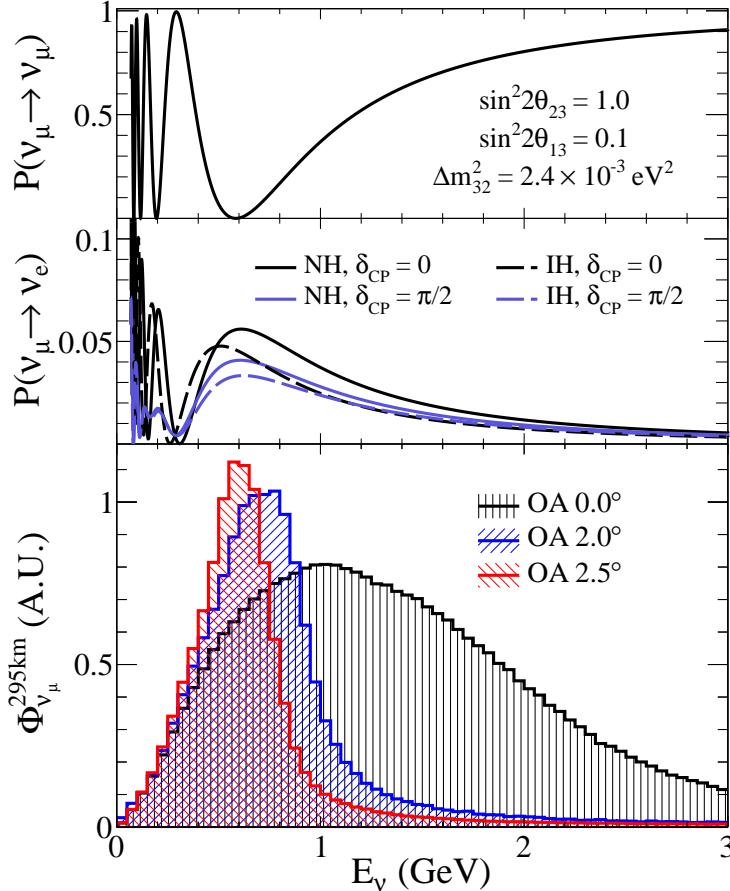
803     The beam dump, situated at the end of the decay volume, stops all charged  
 804     particles other than highly energetic muons ( $p_\mu > 5\text{GeV}$ ). The MuMon detector  
 805     monitors the penetrating muons to determine the beam direction and inten-  
 806     sity which is used to constrain some of the beam flux systematics within the  
 807     analysis [116, 118].

808     The T2K experiment uses an off-axis beam to narrow the neutrino energy  
 809     distribution. This was the first implementation of this technique in a long-  
 810     baseline neutrino oscillation experiment after its original proposal [119]. Pion  
 811     decay,  $\pi \rightarrow \mu + \nu_\mu$ , is a two-body decay. Consequently, the neutrino energy,  
 812      $E_\nu$ , can be determined based on the pion energy,  $E_\pi$ , and the angle at which  
 813     the neutrino is emitted,  $\theta$ ,

$$E_\nu = \frac{m_\pi^2 - m_\mu^2}{2(E_\pi - p_\pi \cos(\theta))}, \quad (3.5)$$

814     where  $m_\pi$  and  $m_\mu$  are the mass of the pion and muon respectively. For a fixed  
 815     energy pion, the neutrino energy distribution is dependent upon the angle at  
 816     which the neutrinos are observed from the initial pion beam direction. For the  
 817     295km baseline at T2K,  $E_\nu = 0.6\text{GeV}$  maximises the electron neutrino appearance  
 818     probability,  $P(\nu_\mu \rightarrow \nu_e)$ , whilst minimising the muon disappearance probability,

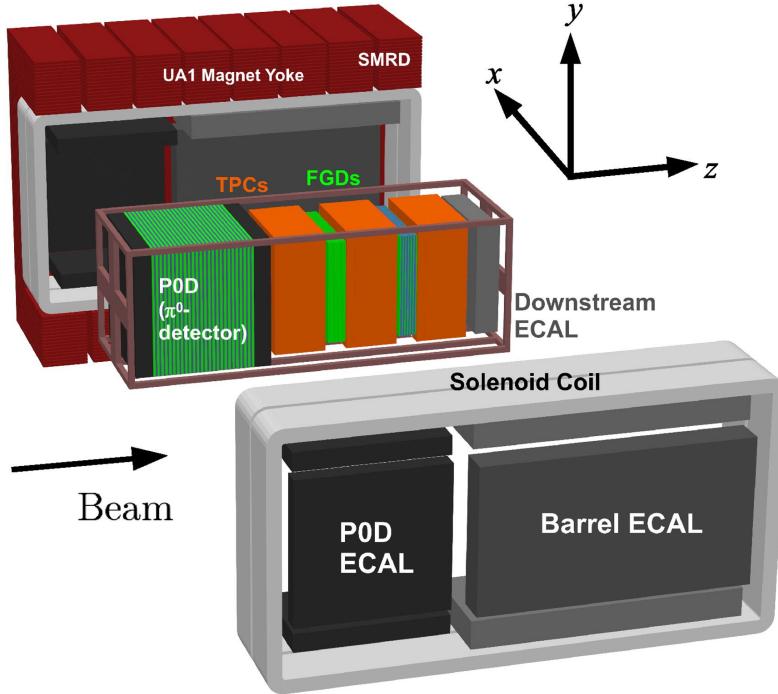
<sup>819</sup>  $P(\nu_\mu \rightarrow \nu_\mu)$ . Figure 3.8 illustrates the neutrino energy distribution for a range of  
<sup>820</sup> off-axis angles, as well as the oscillation probabilities most relevant to T2K.



**Figure 3.8:** Top panel: T2K muon neutrino disappearance probability as a function of neutrino energy. Middle panel: T2K electron neutrino appearance probability as a function of neutrino energy. Bottom panel: The neutrino flux distribution for three different off-axis angles (Arbitrary units) as a function of neutrino energy.

### 3.2.3 The Near Detector at 280m

<sup>821</sup> Whilst all the near detectors are situated in the same “pit” located at 280m from  
<sup>822</sup> the beamline, the “ND280” detector is the off-axis detector which is situated at  
<sup>823</sup> the same off-axis angle as the Super-Kamiokande far detector. It has two primary  
<sup>824</sup> functions; firstly it measures the neutrino flux and secondly, it counts the event  
<sup>825</sup> rates of different types of neutrino interactions. Both of these constrain the flux  
<sup>826</sup> and cross-section systematics invoked within the model for a more accurate  
<sup>827</sup> prediction of the expected event rate at the far detector.  
<sup>828</sup>



**Figure 3.9:** The components of the ND280 detector. The neutrino beam travels from left to right. Taken from [115].

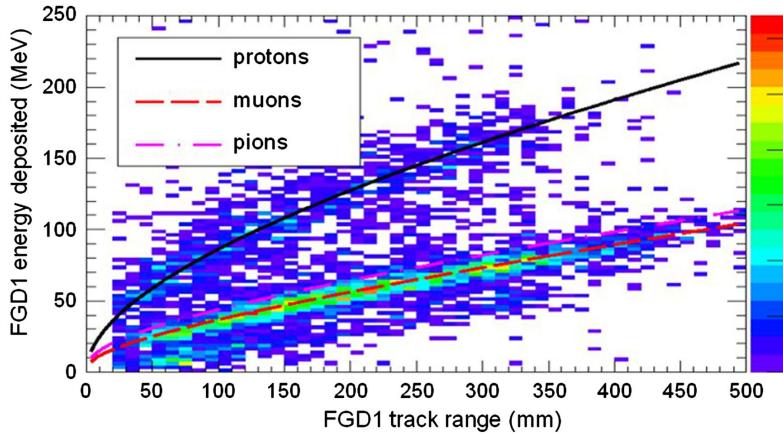
As illustrated in Figure 3.9, the ND280 detector consists of several sub-detectors. The most important part of the detector for this analysis is the tracker region. This is comprised of two-time projection chambers (TPCs) sandwiched between three fine grain detectors (FGDs). The FGDs contain both hydrocarbon plastics and water targets for neutrino interactions and provide track reconstruction near the interaction vertex. The emitted charged particles can then propagate into the TPCs which provide particle identification and momentum reconstruction. The FGDs and TPCs are further described in subsubsection 3.2.3.1 and subsubsection 3.2.3.2 respectively. The electromagnetic calorimeter (ECAL) encapsulates the tracker region alongside the  $\pi^0$  detector (P0D). The ECAL measures the deposited energy from photons emitted from interactions within the FGD. The P0D constrains the cross-section of neutral current interactions which generate neutral pions, which is one of the largest backgrounds in the electron neutrino appearance oscillation channel. The P0D and ECAL detectors are detailed in subsubsection 3.2.3.3 and subsubsection 3.2.3.4 respectively. The entire detector is located within a large yoke magnet which produces a 0.2T magnetic field.

845 field. This design of the magnet also includes a scintillating detector called the  
846 side muon range detector (SMRD), which is used to track high-angle muons as  
847 well as acting as a cosmic veto. The SMRD is described in subsubsection 3.2.3.5.

848 **3.2.3.1 Fine Grained Detectors**

849 The T2K tracker region is comprised of two fine-grained detectors (FGD) and  
850 three Time Projection Chambers (TPC). A detailed description of the FGD design,  
851 construction, and assembly is found in [120] and summarised below. The FGDS  
852 are the primary target for neutrino interactions with a mass of 1.1 tonnes per FGD.  
853 Alongside this, the FGDS are designed to be able to track short-range particles  
854 which do not exit the FGD. Typically, short-range particles are low momentum  
855 and are observed as tracks that deposit a large amount of energy per unit length.  
856 This means the FGD needs good granularity to resolve these particles. The  
857 FGDS have the best timing resolution ( $\sim 3\text{ns}$ ) of any of the sub-detectors of the  
858 ND280 detector. As such, the FGDS are used for time of flight measurements  
859 to distinguish forward-going positively charged particles from backward-going  
860 negatively charged particles. Finally, any tracks which pass through multiple  
861 sub-detectors are required to be track matched to the FGD.

862 Both FGDS are made from square scintillator planes of side length 186cm and  
863 width 2.02cm. Each plane consists of two layers of 192 scintillator bars in an  
864 X or Y orientation. A wavelength-shifting fiber is threaded through the center  
865 of each bar and is read out by a multi-pixel photon counter (MPPC). FGD1 is  
866 the most upstream of the two FGDS and contains 15 planes of carbon plastic  
867 scintillator which is a common target in external neutrino scattering data. As  
868 the far detector is a pure water target, 7 of the 15 scintillator planes in FGD2  
869 have been replaced with a hybrid water-scintillator target. Due to the complexity  
870 of the nucleus, nuclear effects can not be extrapolated between different nuclei.  
871 Therefore having the ability to take data on one target which is the same as  
872 external data and another target which is the same as the far detector target is  
873 beneficial for reliable model parameter estimates.



**Figure 3.10:** Comparison of data to Monte Carlo prediction of integrated deposited energy as a function of track length for particles that stopped in FGD1. Taken from [120].

874     The integrated deposited energy is used for particle identification. The FGD  
875     can distinguish protons from other charged particles by comparing the integrated  
876     deposited energy from data to Monte Carlo prediction as seen in Figure 3.10.

### 877     3.2.3.2 Time Projection Chambers

878     The majority of particle identification and momentum measurements within  
879     ND280 are provided by three Time Projection Chambers (TPCs) [121]. The  
880     TPCs are located on either side of the FGDs. They are located inside of the  
881     magnetic field meaning the momentum of a charged particle can be determined  
882     from the bending of the track.

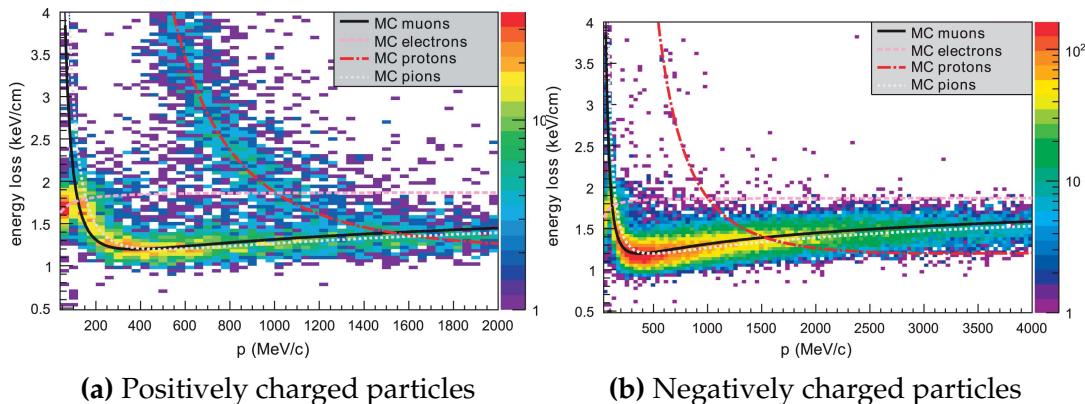
883     Each TPC module consists of two gas-tight boxes, as shown in Figure 3.11,  
884     which are made of non-magnetic material. The outer box is filled with CO<sub>2</sub> which  
885     acts as an electrical insulator between the inner box and the ground. The inner box  
886     forms the field cage which produces a uniform electric drift field of  $\sim 275\text{V/cm}$   
887     and is filled with an argon gas mixture. Charged particles moving through this  
888     gas mixture ionize the gas and the ionised charge is drifted towards micromegas  
889     detectors which measure the ionization charge. The time and position information  
890     in the readout allows a three-dimensional image of the neutrino interaction.

891     The particle identification of tracks that pass through the TPCs is performed  
892     using dE/dx measurements. Figure 3.12 illustrates the data to Monte Carlo



**Figure 3.11:** Schematic design of a Time Projection Chamber detector. Taken from [121].

893 distributions of the energy lost by a charged particle passing through the TPC as  
 894 a function of the reconstructed particle momentum. The resolution is  $7.8 \pm 0.2\%$   
 895 meaning that electrons and muons can be distinguished. This allows reliable  
 896 measurements of the intrinsic  $\nu_e$  component of the beam.



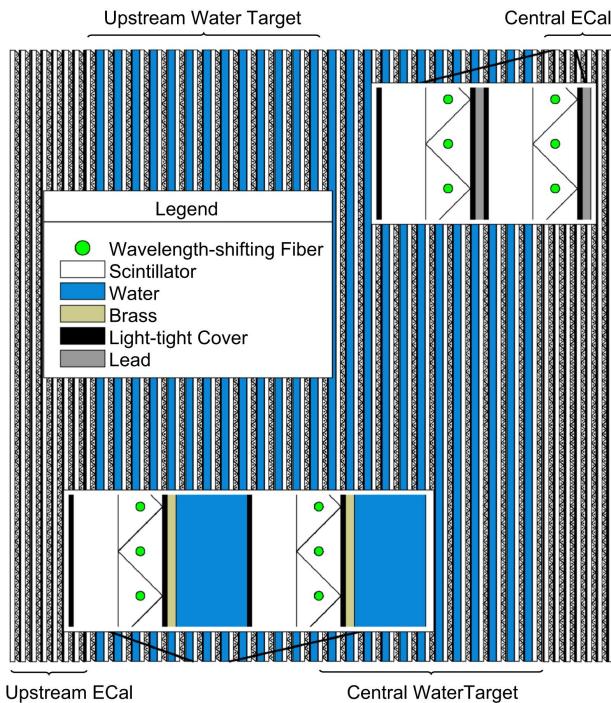
**Figure 3.12:** The distribution of energy loss as a function of reconstructed momentum for charged particles passing through the TPC, comparing data to Monte Carlo prediction. Taken from [121].

897 **3.2.3.3  $\pi^0$  Detector**

898 If one of the  $\gamma$ -rays from a  $\pi^0 \rightarrow 2\gamma$  decay is missed at the far detector, the  
 899 reconstruction will determine that event to be a charge current  $\nu_e$ -like event.  
 900 This is one of the main backgrounds hindering the electron neutrino appearance

901 searches. The  $\pi^0$  detector (P0D) measures the cross-section of the neutral current  
 902 induced neutral pion production on a water target to constrain this background.

903 The P0D is a cube of approximately 2.5m length consisting of layers of scin-  
 904 tillating bars, brass and lead sheets, and water bags as illustrated in Figure 3.13.  
 905 Two electromagnetic calorimeters are positioned at the most upstream and most  
 906 downstream position in the sub-detector and the water target is situated in  
 907 between them. The scintillator layers are built from two triangular bars orientated  
 908 in opposite directions to form a rectangular layer. Each triangular scintillator bar  
 909 is threaded with optical fiber which is read out by MPPCs. The high-Z brass and  
 910 lead regions produce electron showers from the photons emitted in  $\pi^0$  decay.



**Figure 3.13:** A schematic of the P0D side-view. Taken from [122].

911 The sub-detector can generate measurements of NC1 $\pi^0$  cross-sections on a  
 912 water target by measuring the event rate both with and without the water target,  
 913 with the cross-section on a water target being determined as the difference. The to-  
 914 tal active mass is 16.1 tonnes when filled with water and 13.3 tonnes when empty.

915 **3.2.3.4 Electromagnetic Calorimeter**

916 The electromagnetic calorimeter [123] (ECal) encapsulates the P0D and tracking  
 917 sub-detectors. Its primary purpose is to aid  $\pi^0$  reconstruction from any interac-  
 918 tion in the tracker. To do this, it measures the energy and direction of photon  
 919 showers from  $\pi^0 \rightarrow 2\gamma$  decay. It can also distinguish pion and muon tracks  
 920 depending on the shape of the photon shower deposited.

921 The ECal is comprised of three sections; the P0D ECal which surrounds the  
 922 P0D, the barrel ECal which encompasses the tracking region, and the downstream  
 923 ECal which is situated downstream of the tracker region. The barrel and down-  
 924 stream ECals are tracking calorimeters that focus on electromagnetic showers  
 925 from high-angle particles emitted from the tracking sub-detectors. Particularly in  
 926 the TPC, high-angle tracks (those which travel perpendicularly to the beam-axis)  
 927 can travel along a single scintillator bar resulting in very few hits. The width of  
 928 the barrel and downstream ECal corresponds to  $\sim 11$  electron radiation lengths  
 929 to ensure a significant amount of the  $\pi^0$  energy is contained. As the P0D has  
 930 its own calorimetry which reconstructs showers, the P0D ECal determines the  
 931 energy which escapes the P0D.

932 Each ECal is constructed of multiple layers of scintillating bars sandwiched  
 933 between lead sheets. The scintillating bars are threaded with optical fiber and read  
 934 out by MPPCs. Each sequential layer of the scintillator is orientated perpendicular  
 935 to the previous which allows a three-dimensional event reconstruction. The  
 936 target mass of the P0D ECal, barrel ECal, and downstream ECal are 1.50, 4.80,  
 937 and 6.62 tonnes respectively.

938 **3.2.3.5 Side Muon Range Detector**

939 As illustrated in Figure 3.9, the ECal, FGDs, P0D, and TPCs are enclosed within  
 940 the UA1 magnet. Reconditioned after use in the UA1 [124] and NOMAD [125]  
 941 experiments, this magnet provides a uniform horizontal magnetic field of 0.2T  
 942 with an uncertainty of  $2 \times 10^{-4}$ T.

943     Built into the UA1 magnet, the side muon range detector (SMRD)[126] monitors  
944     high-energy muons which leave the tracking region and permeate through  
945     the ECal. It additionally acts as a cosmic muon veto and trigger.

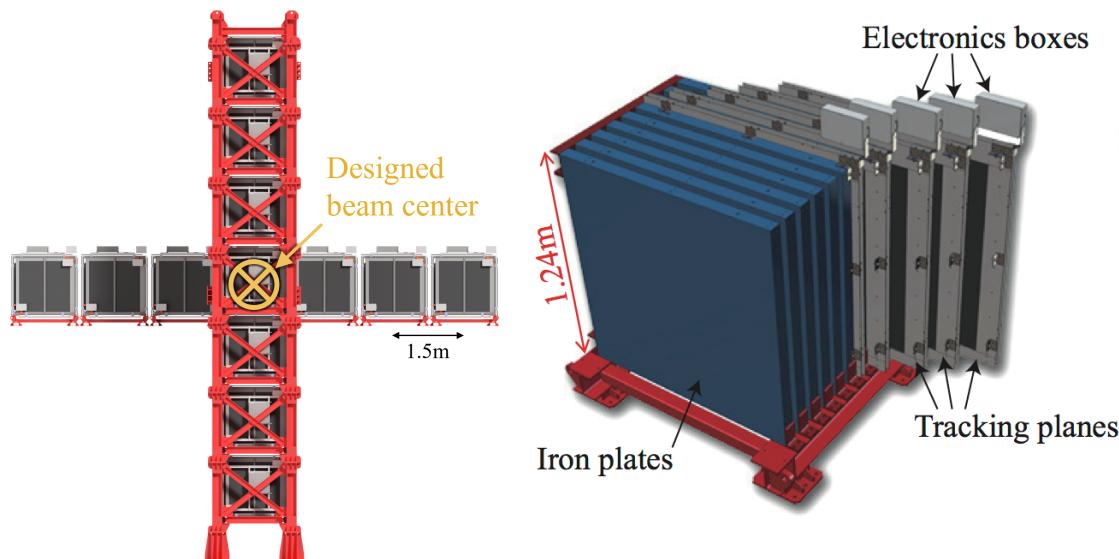
#### 946     **3.2.4 The Interactive Neutrino GRID**

947     The Interactive Neutrino GRID (INGRID) detector is situated within the same  
948     “pit” as the other near detectors. It is aligned with the beam in the “on-axis”  
949     position and measures the beam direction, spread, and intensity. The detector  
950     was originally designed with 16 identical modules [115] (two modules have since  
951     been decommissioned) and a “proton” module. The design of the detector is 14  
952     modules oriented in a cross with length and height  $10\text{m} \times 10\text{m}$ , as illustrated  
953     in Figure 3.14.

954     Each module is composed of iron sheets interlaced with eleven tracking  
955     scintillator planes for a total target mass of 7.1 tonnes per module. The scintillator  
956     design is an X-Y pattern of 24 bars in both orientations, where each bar contains  
957     wave-length shifting fibers which are connected to multi-pixel photon counters  
958     (MPPCs). Each module is encapsulated inside veto planes to aid the rejection  
959     of charged particles entering the module.

960     The proton module is different from the other modules in that it consists  
961     of entirely scintillator planes with no iron target. The scintillator bars are also  
962     smaller than those used in the other modules to increase the granularity of  
963     the detector and improve tracking capabilities. The module sits in the center  
964     of the beamline and is designed to give precise measurements of quasi-elastic  
965     charged current interactions to evaluate the performance of the Monte Carlo  
966     simulation of the beamline.

967     The INGRID detector can measure the beam direction to an uncertainty of  
968     0.4mrad and the beam centre within a resolution of 10cm [115]. The beam  
969     direction in both the vertical and horizontal directions is discussed in [127]  
970     and it is found to be in good agreement with the MUMON monitor described  
971     in subsection 3.2.2.



**Figure 3.14:** Left panel: The Interactive Neutrino GRID on-axis Detector. 14 modules are arranged in a cross-shape configuration, with the center modules being directly aligned with the on-axis beam. Right panel: The layout of a single module of the INGRID detector. Both figures are recreated from [115].

# 4

972

973

974

## Bayesian Statistics and Markov Chain Monte Carlo Techniques

975 This thesis presents a Bayesian oscillation analysis. To extract the oscillation  
976 parameters, a Markov Chain Monte Carlo (MCMC) method is used. This chapter  
977 explains the theory of how parameter estimates can be determined using this  
978 technique and condenses the material found in the literature [128–131].

979 The oscillation parameter determination presented here is built upon a si-  
980 multaneous fit to neutrino beam data in the near detector, beam data at SK, and  
981 atmospheric data at SK. In total, there are four oscillation parameters of interest  
982 ( $\sin^2(\theta_{23})$ ,  $\sin^2(\theta_{13})$ ,  $\Delta m_{32}^2$ , and  $\delta_{CP}$ ), two oscillation parameters to which this  
983 study will not be sensitive ( $\sin^2(\theta_{12})$ ,  $\Delta m_{21}^2$ ) and many nuisance parameters that  
984 control the systematic uncertainty models.

985 This analysis uses a Monte Carlo technique to generate a multi-dimensional  
986 probability distribution across all of the model parameters used in the fit. To  
987 determine an estimate for each parameter, this multi-dimensional object is in-  
988 tegrated over all other parameters. This process is called Marginalisation and  
989 is described in subsection 4.3.1. Monte Carlo techniques approximate the prob-  
990 ability distribution of each parameter within the limit of generating infinite  
991 samples. As ever, generating a large number of samples is time and resource-

dependent. Therefore, an MCMC technique is utilised within this analysis to reduce the required number of steps to sufficiently sample the parameter space.

This technique is described in further detail in subsection 4.2.1.

DB: Introduce MaCh3 and say what I did on it

## 4.1 Bayesian Statistics

Bayesian inference treats observable data,  $D$ , and model parameters,  $\vec{\theta}$ , on equal footing such that a probability model of both data and parameters is required. This is the joint probability distribution  $P(D, \vec{\theta})$  and can be described by the prior distribution for model parameters  $P(\vec{\theta})$  and the likelihood of the data given the model parameters  $P(D|\vec{\theta})$ ,

$$P(D, \vec{\theta}) = P(D|\vec{\theta})P(\vec{\theta}). \quad (4.1)$$

The prior distribution,  $P(\vec{\theta})$ , describes all previous knowledge about the parameters within the model. For example, if the risk of developing health problems is known to increase with age, the prior distribution would describe the increase. For the purpose of this analysis, the prior distribution is typically the best-fit values taken from external data measurements with a Gaussian uncertainty. The prior distribution can also contain correlations between model parameters. In an analysis using Monte Carlo techniques, the likelihood of measuring some data assuming some set of model parameters is calculated by comparing the Monte Carlo prediction generated at that particular set of model parameters to the data.

It is parameter estimation that is important for this analysis and as such, Bayes' theorem [132] is applied to calculate the probability for each parameter to have a certain value given the observed data,  $P(\vec{\theta}|D)$ , which is known as the posterior distribution (often termed the posterior). This can be expressed as

$$P(\vec{\theta}|D) = \frac{P(D|\vec{\theta})P(\vec{\theta})}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}}. \quad (4.2)$$

1016        The denominator in Equation 4.2 is the integral of the joint probability distri-  
 1017        bution over all values of all parameters used within the fit. For brevity, the  
 1018        posterior distribution is

$$P(\vec{\theta}|D) \propto P(D|\vec{\theta})P(\vec{\theta}). \quad (4.3)$$

1019        For the purposes of this analysis, it is acceptable to neglect the normalisation  
 1020        term and focus on this proportional relationship.

### 1021        4.1.1 Application of Prior Knowledge

1022        The posterior distribution is proportional to the prior uncertainty applied to  
 1023        each parameter, as illustrated by Equation 4.3. This means that it is possible  
 1024        to change the prior after the posterior distribution has been determined. The  
 1025        prior uncertainty of a particular parameter can be ‘divided’ out of the posterior  
 1026        distribution and the resulting distribution can be reweighted using the new  
 1027        prior uncertainty that is to be applied. The methodology and implementation  
 1028        of changing the prior follows that described in [133].

1029        An example implementation that is useful for this analysis is the application  
 1030        of the “reactor constraint”. As discussed in section 2.4, an external constraint  
 1031        on  $\sin^2(\theta_{13})$  is determined from measurements taken from reactor experiments.  
 1032        However, the sensitivities from just using the T2K and SK samples is equally  
 1033        as important. Without this technique, two fits would have to be run, doubling  
 1034        the required resources. Therefore, the key benefit for this analysis is the fact that  
 1035        only a single ‘fit’ has to be performed and can be used to build the two posterior  
 1036        distributions of the with and without reactor constraint applied.

## 1037        4.2 Monte Carlo Simulation

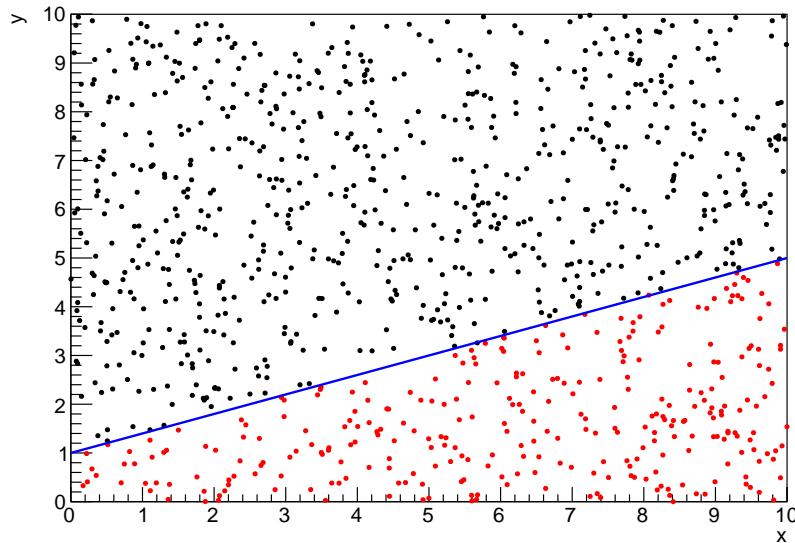
1038        Monte Carlo techniques are used to numerically solve a complex problem that  
 1039        does not necessarily have an analytical solution. These techniques rely on

1040 building a large ensemble of samples from an unknown distribution and then  
1041 using the ensemble to approximate the properties of the distribution.

1042 An example that uses Monte Carlo techniques is to calculate the area under-  
1043 neath a curve. For example, take the problem of calculating the area under a  
1044 straight line with gradient  $M = 0.4$  and intercept  $C = 1.0$ . Analytically, one can  
1045 calculate the area under the line is equal to 30 units for  $0 \leq x \leq 10$ . Using Monte  
1046 Carlo techniques, one can calculate the area under this line by throwing many  
1047 random values for the  $x$  and  $y$  components of each sample and then calculating  
1048 whether that point falls below the line. The area can then be calculated by the  
1049 ratio of points below the line to the total number of samples thrown multiplied by  
1050 the total area in which samples were scattered. The study is shown in Figure 4.1  
1051 highlights this technique and finds the area under the curve to be 29.9 compared  
1052 to an analytical solution of 30.0. The deviation of the numerical to analytical  
1053 solution can be attributed to the number of samples used in the study. The  
1054 accuracy of the approximation in which the properties of the Monte Carlo samples  
1055 replicate those of the desired distribution is dependent on the number of samples  
1056 used. Replicating this study with a differing number of Monte Carlo samples  
1057 used in each study (As shown in Figure 4.2) highlights how the Monte Carlo  
1058 techniques are only accurate within the limit of a high number of samples.

1059 Whilst the above example has an analytical solution, these techniques are just  
1060 as applicable to complex solutions. Clearly, any numerical solution is only as  
1061 useful as its efficiency. As discussed, the accuracy of the Monte Carlo technique is  
1062 dependent upon the number of samples generated to approximate the properties  
1063 of the distribution. Furthermore, if the positions at which the samples are  
1064 evaluated are not ‘cleverly’ picked, the efficiency of the Monte Carlo technique  
1065 significantly drops. Given the example in Figure 4.1, if the region in which the  
1066 samples are scattered significantly extends passed the region of interest, many  
1067 calculations will be calculated but do not add to the ability of the Monte Carlo  
1068 technique to achieve the correct result. For instance, any sample evaluated at  
1069 a  $y \geq 5$  could be removed without affecting the final result. This does bring in

- 1070 an aspect of the ‘chicken and egg’ problem in that to achieve efficient sampling,  
1071 one needs to know the distribution beforehand.



**Figure 4.1:** Example of using Monte Carlo techniques to find the area under the blue line. The gradient and intercept of the line are 0.4 and 1.0 respectively. The area found to be under the curve using one thousand samples is 29.9 units.

### 1072 4.2.1 Markov Chain Monte Carlo

- 1073 This analysis utilises a multi-dimensional probability distribution, with some  
1074 dimensions being significantly more constrained than others. These constraints  
1075 can be from prior knowledge of parameter distributions from external data or  
1076 un-physical regions in which parameters can not exist. To maximise the efficiency  
1077 of building the posterior distribution, a Markov Chain Monte Carlo (MCMC)  
1078 technique is used. This employs a Markov chain to select the points at which  
1079 to sample the posterior distribution. It performs a semi-random stochastic walk  
1080 through the allowable parameter space. This builds a posterior distribution  
1081 which has the property that the density of sampled points is proportional to the  
1082 probability density of that parameter. This means that the samples produced by  
1083 this technique are not statistically independent but they will cover the space  
1084 of the distribution.



**Figure 4.2:** The area under a line of gradient 0.4 and intercept 1.0 for the range  $0 \leq x \leq 10$  as calculated using Monte Carlo techniques as a function of the number of samples used in each repetition. The analytical solution to the area is 30 units as given by the red line.

1085     A Markov chain functions by selecting the position of step  $\vec{x}_{i+1}$  based on the  
 1086     position of  $\vec{x}_i$ . The space in which the Markov chain selects samples is dependent  
 1087     upon the total number of parameters utilised within the fit, where a discrete point  
 1088     in this space is described by the N-dimensional space  $\vec{x}$ . In a perfectly operating  
 1089     Markov chain, the position of the next step depends solely on the previous step  
 1090     and not on the further history of the chain ( $\vec{x}_0, \vec{x}_1$ , etc.). However, in solving  
 1091     the multi-dimensionality of the fit used within this analysis, each step becomes  
 1092     correlated with several of the steps preceding itself. Providing the MCMC chain is  
 1093     well optimised, it will begin to converge towards a unique stationary distribution.  
 1094     The period between the chain's initial starting point and the convergence to the  
 1095     unique stationary distribution is colloquially known as the burn-in period. Once  
 1096     the chain reaches the stationary distribution, all points sampled after that point  
 1097     will look like samples from that distribution.

1098     Further details of the theories underpinning MCMC techniques are discussed  
 1099     in [129] but can be summarised by the requirement that the chain satisfies the  
 1100     three 'regularity conditions':

- 1101     • Irreducibility: From every position in the parameter space  $\vec{x}$ , there must  
1102       exist a non-zero probability for every other position in the parameter space  
1103       to be reached.
- 1104     • Recurrence: Once the chain arrives at the stationary distribution, every step  
1105       following from that position must be samples from the same stationary  
1106       distribution.
- 1107     • Aperiodicity: The chain must not repeat the same sequence of steps at any  
1108       point throughout the sampling period.

1109     The output of the chain after burn-in (i.e. the sampled points after the chain  
1110    has reached the stationary distribution) can be used to approximate the posterior  
1111    distribution and model parameters  $\vec{\theta}$ . To achieve the requirement that the unique  
1112    stationary distribution found by the chain be the posterior distribution, one  
1113    can use the Metropolis-Hastings algorithm. This guides the stochastic process  
1114    depending on the likelihood of the current proposed step compared to that  
1115    of the previous step.

### 1116    **4.2.2 Metropolis-Hastings Algorithm**

1117    As a requirement for MCMCs, the Markov chain implemented in this technique  
1118    must have a unique stationary distribution that is equivalent to the posterior  
1119    distribution. To ensure this requirement and that the regularity conditions are  
1120    met, this analysis utilises the Metropolis-Hastings (MH) algorithm [134, 135].  
1121    For the  $i^{th}$  step in the chain, the MH algorithm determines the position in the  
1122    parameter space to which the chain moves to based on the current step,  $\vec{x}_i$ , and  
1123    the proposed step,  $\vec{y}_{i+1}$ . The proposed step is randomly selected from some  
1124    proposal function  $f(\vec{x}_{i+1}|\vec{x}_i)$ , which depends solely on the current step (ie. not  
1125    the further history of the chain). The next step in the chain  $\vec{x}_{i+1}$  can be either the  
1126    current step or the proposed step determined by whether the proposed step is

<sub>1127</sub> accepted or rejected. To decide if the proposed step is selected, the acceptance  
<sub>1128</sub> probability,  $\alpha(\vec{x}_i, \vec{y}_i)$ , is calculated as

$$\alpha(\vec{x}_i, \vec{y}_i) = \min \left( 1, \frac{P(\vec{y}_{i+1}|D)f(\vec{x}_i|\vec{y}_{i+1})}{P(\vec{x}_i|D)f(\vec{y}_{i+1}|\vec{x}_i)} \right). \quad (4.4)$$

<sub>1129</sub> Where  $P(\vec{y}_{i+1}|D)$  is the posterior distribution as introduced in section 4.1. To  
<sub>1130</sub> simplify this calculation, the proposal function is required to be symmetric such  
<sub>1131</sub> that  $f(\vec{x}_i|\vec{y}_{i+1}) = f(\vec{y}_{i+1}|\vec{x}_i)$ . In practice, a multi-variate Gaussian distribution is  
<sub>1132</sub> used to throw parameter proposals. This reduces Equation 4.4 to

$$\alpha(\vec{x}_i, \vec{y}_i) = \min \left( 1, \frac{P(\vec{y}_{i+1}|D)}{P(\vec{x}_i|D)} \right). \quad (4.5)$$

<sub>1133</sub> **DB: Figure out what Giles means**

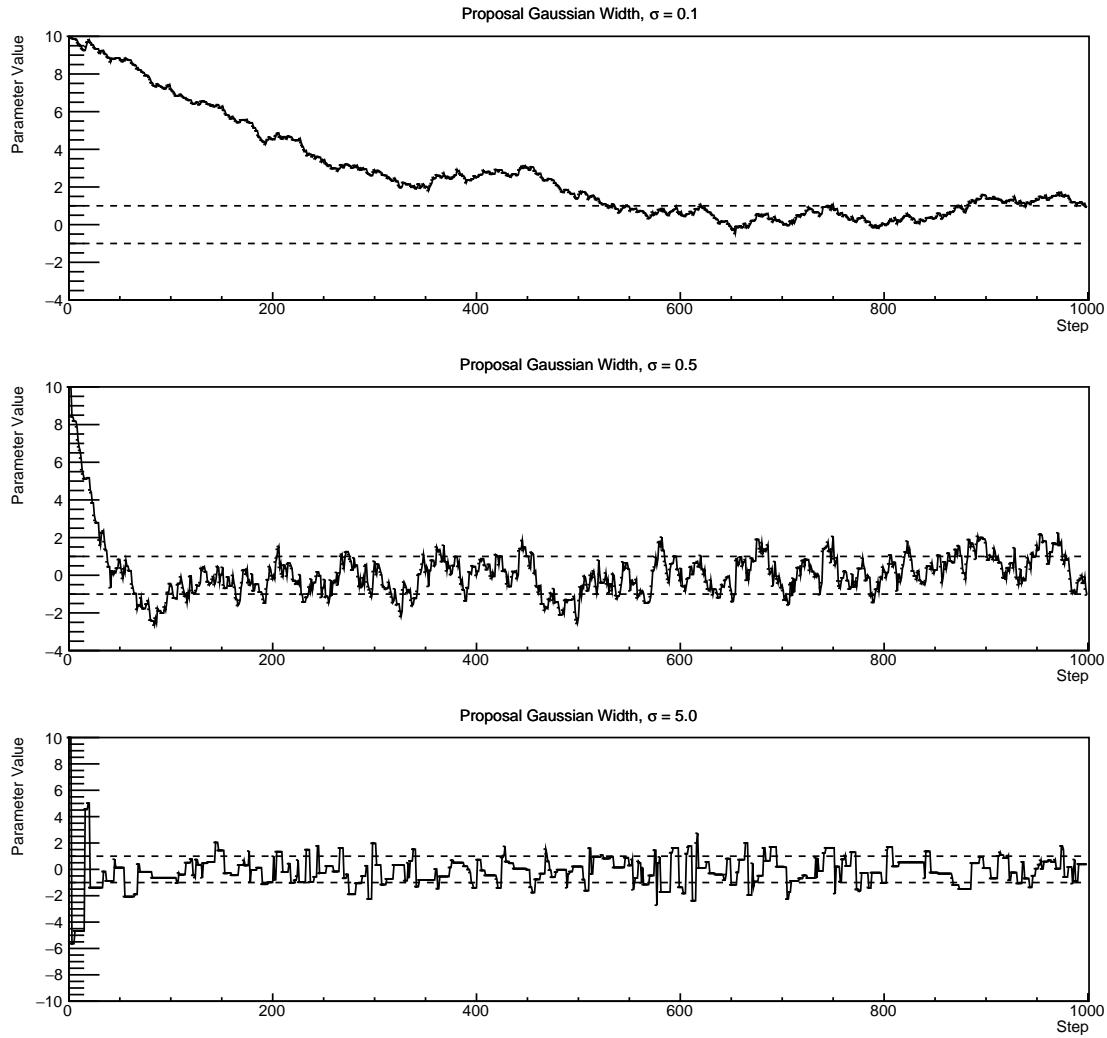
<sub>1134</sub> After calculating this quantity, a random number,  $\beta$ , is generated uniformly  
<sub>1135</sub> between 0 and 1. If  $\beta \leq \alpha(\vec{x}_i, \vec{y}_i)$ , the proposed step is accepted. Otherwise,  
<sub>1136</sub> the chain sets the next step equal to the current step. This procedure is repeated  
<sub>1137</sub> for subsequent steps. This can be interpreted as if the posterior probability  
<sub>1138</sub> of the proposed step is greater than that of the current step, ( $P(\vec{y}_{i+1}|D) \geq$   
<sub>1139</sub>  $P(\vec{x}_i|D)$ ), the proposed step will always be accepted. If the opposite is true,  
<sub>1140</sub> ( $P(\vec{y}_{i+1}|D) \leq P(\vec{x}_i|D)$ ), the proposed step will be accepted with probability  
<sub>1141</sub>  $P(\vec{x}_i|D)/P(\vec{y}_{i+1}|D)$ . This ensures that the Markov chain does not get trapped  
<sub>1142</sub> in any local minima in the potentially non-Gaussian posterior distribution. The  
<sub>1143</sub> outcome of this technique is that the density of steps taken in a discrete region  
<sub>1144</sub> is directly proportional to the probability density in that region.

### <sub>1145</sub> 4.2.3 MCMC Optimisation

<sub>1146</sub> As discussed in subsection 4.2.2, the proposal function invoked within the MH  
<sub>1147</sub> algorithm can take any form and the chain will still converge to the stationary  
<sub>1148</sub> distribution. At each set of proposed parameter values, a prediction of the same  
<sub>1149</sub> spectra has to be generated which requires significant computational resources.  
<sub>1150</sub> Therefore, the number of steps taken before the unique stationary distribution  
<sub>1151</sub> is found should be minimised as only steps after convergence add information

1152 to the oscillation analysis. Furthermore, the chain should entirely cover the  
1153 allowable parameter space to ensure that all values have been considered. Tuning  
1154 the distance that the proposal function jumps between steps on a parameter-by-  
1155 parameter basis can both minimise the length of the burn-in period and ensure  
1156 that the correlation between step  $\vec{x}_i$  and  $\vec{x}_j$  is sufficiently small.

1157 The effect of changing the width of the proposal function is highlighted in  
1158 Figure 4.3. Three scenarios, each with the same underlying stationary distribution  
1159 (A Gaussian of width 1.0 and mean 0.), are presented. The only difference between  
1160 the three scenarios is the width of the proposal function, colloquially known as  
1161 the ‘step size  $\sigma$ ’. Each scenario starts at an initial parameter value of 10.0 which  
1162 would be considered an extreme variation. For the case where  $\sigma = 0.1$ , it is  
1163 clear to see that the chain takes a long time to reach the expected region of the  
1164 parameter. This indicates that this chain would have a large burn-in period and  
1165 does not converge to the stationary distribution until step  $\sim 500$ . Furthermore,  
1166 whilst the chain does move towards the expected region, each step is significantly  
1167 correlated with the previous. Considering the case where  $\sigma = 5.0$ , the chain  
1168 approaches the expected parameter region almost instantly meaning that the  
1169 burn-in period is not significant. However, there are clearly large regions of steps  
1170 where the chain does not move. This is likely due to the chain proposing steps  
1171 in the tails of the distribution which have a low probability of being accepted.  
1172 Consequently, this chain would take a significant number of steps to fully span  
1173 the allowable parameter region. For the final scenario, where  $\sigma = 0.5$ , you can  
1174 see a relatively small burn-in period of approximately 100 steps. Once the chain  
1175 reaches the stationary distribution, it moves throughout the expected region of  
1176 parameter values many times, sufficiently sampling the full parameter region.  
1177 This example is a single parameter varying across a continuous distribution and  
1178 does not fully reflect the difficulties in the many-hundred multi-variate parameter  
1179 distribution used within this analysis. However, it does give a conceptual idea of  
1180 the importance of selecting the proposal function and associated step size.



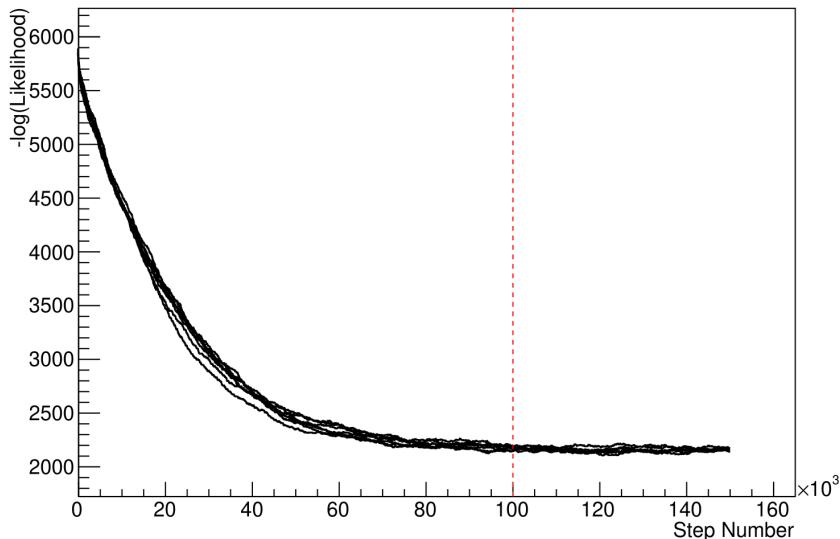
**Figure 4.3:** Three MCMC chains, each with a stationary distribution equal to a Gaussian centered at 0 and width 1 (As indicated by the black dotted lines). All of the chains use a Gaussian proposal function but have different widths (or ‘step size  $\sigma$ ’). The top panel has  $\sigma = 0.1$ , middle panel has  $\sigma = 0.5$  and the bottom panel has  $\sigma = 5.0$ .

1181 As discussed, step size tuning directly correlates to the average step accep-  
 1182 tance rate. If the step size is too small, many steps will be accepted but the  
 1183 chain moves slowly. If the opposite is true, many steps will be rejected as the  
 1184 chain proposes steps in the tails of the distribution. Discussion in [136] suggests  
 1185 that the ‘ideal’ acceptance rate of a high dimension MCMC chain should be  
 1186 approximately  $\sim 25\%$ . An “ideal” step size [136] of

$$\sigma = \frac{2.4}{N_p}, \quad (4.6)$$

1187 where  $N_p$  is the number of parameters included in the MCMC fit. However,  
 1188 the complex correlations between systematics mean that some parameters have  
 1189 to be hand-tuned and many efforts have been taken to select a set of parameter-  
 1190 by-parameter step sizes to approximately reach the ideal acceptance rate.

1191 Figure 4.4 highlights the likelihood as calculated by the fit in DB: [Link to](#)  
 1192 [AsimovA Sensitivity Section](#) as a function of the number of steps in each chain. In  
 1193 practice, many independent MCMC chains are run simultaneously to parallelise  
 1194 the task of performing the fit. This figure overlays the distribution found in each  
 1195 chain. As seen, the likelihood decreases from its initial value and converges  
 1196 towards a stationary distribution after  $\sim 1 \times 10^5$  steps.



**Figure 4.4:** The log-likelihood from the fit detailed in DB: [Link to](#) [AsimovA Sensitivity Section](#) as a function of the number of steps accumulated in each fit. Many independent MCMC chains were run in parallel and overlaid on this plot. The red line indicates the  $1 \times 10^5$  step burn-in period after which the log-likelihood becomes stable.

1197 Multiple configurations of this analysis have been performed throughout this  
 1198 thesis where different samples or systematics have been used. For all of these con-  
 1199 figurations, it was found that a burnin period of  $1 \times 10^5$  was sufficient in all cases.

## 1200 4.3 Understanding the MCMC Results

1201 The previous sections have described how to generate the posterior probability  
1202 distribution using Bayesian MCMC techniques. However, this analysis focuses  
1203 on oscillation parameter determination. The posterior distribution output from  
1204 the chain is a high-dimension object, with as many dimensions as there are  
1205 parameters included in the oscillation analysis. However, this multi-dimensional  
1206 object is difficult to conceptualize so parameter estimations are often presented  
1207 in one or two-dimensional projections of this probability distribution. To do  
1208 this, marginalisation techniques are invoked.

### 1209 4.3.1 Marginalisation

1210 The output of the MCMC chain is a highly dimensional probability distribution  
1211 which is very difficult to interpret. From the standpoint of an oscillation analysis  
1212 experiment, the one or two-dimensional ‘projections’ of the oscillation parameters  
1213 of interest are most relevant. Despite this, the best fit values and uncertainties on  
1214 the oscillation parameters of interest should correctly encapsulate the correlations  
1215 to the other systematic uncertainties (colloquially called ‘nuisance’ parameters).  
1216 For this joint beam and atmospheric analysis, the oscillation parameters of  
1217 interest are  $\sin^2(\theta_{23})$ ,  $\sin^2(\theta_{13})$ ,  $\Delta m_{32}^2$ , and  $\delta_{CP}$ . All other parameters (includ-  
1218 ing the oscillation parameters this fit is insensitive to) are deemed nuisance  
1219 parameters. To generate these projections, the posterior distribution is integrated  
1220 over all nuisance parameters. This is called marginalisation. This technique  
1221 also explains why it is acceptable to neglect the normalisation constant of the  
1222 posterior distribution, which was discussed in section 4.1.

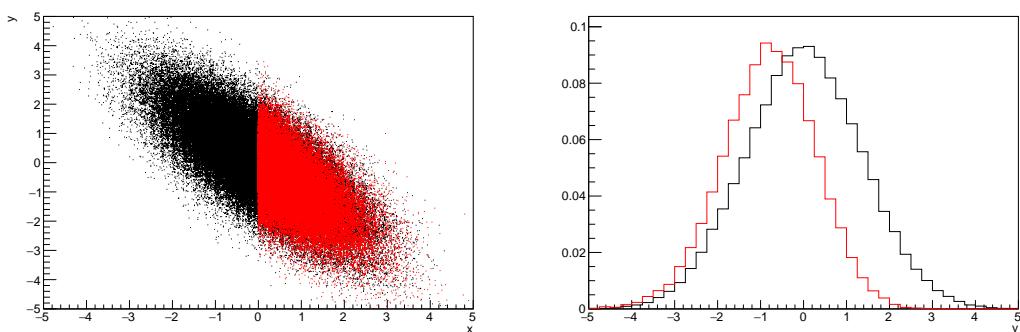
1223 A simple example of the marginalisation technique is to imagine the scenario  
1224 where two coins are flipped. To determine the probability that the first coin  
1225 returned a ‘head’, the exact result of the second coin flip is disregarded and  
1226 simply integrated over. For the parameters of interest,  $\vec{\theta}_i$ , the marginalised

<sub>1227</sub> posterior is calculated by integrating over the nuisance parameters,  $\vec{\theta}_n$ . In this  
<sub>1228</sub> case, Equation 4.2 becomes

$$P(\vec{\theta}_i|D) = \frac{\int P(D|\vec{\theta}_i, \vec{\theta}_n)P(\vec{\theta}_i, \vec{\theta}_n)d\vec{\theta}_n}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}}. \quad (4.7)$$

<sub>1229</sub> Where  $P(\vec{\theta}_i, \vec{\theta}_n)$  encodes the prior knowledge about the uncertainty and  
<sub>1230</sub> correlations between the parameters of interest and the nuisance parameters.  
<sub>1231</sub> In practice, this is simply taking the one or two-dimensional projection of the  
<sub>1232</sub> multi-dimensional probability distribution.

<sub>1233</sub> While in principle an easy solution to a complex problem, correlations be-  
<sub>1234</sub> tween the interesting and nuisance parameters can bias the marginalised results.  
<sub>1235</sub> A similar effect is found when the parameters being marginalised over have  
<sub>1236</sub> non-Gaussian probability distributions. For example, Figure 4.5 highlights the  
<sub>1237</sub> marginalisation bias in the probability distribution found for a parameter when  
<sub>1238</sub> requiring a correlated parameter to have a positive parameter value. Due to  
<sub>1239</sub> the complex nature of the oscillation parameter fit presented in this thesis, there  
<sub>1240</sub> are correlations occurring between the oscillation parameters of interest and the  
<sub>1241</sub> other nuisance parameters included in the fit.



**Figure 4.5:** Left: The two-dimensional probability distribution for two correlated parameters  $x$  and  $y$ . The red distribution shows the two-dimensional probability distribution when  $0 \leq x \leq 5$ . Right: The marginalised probability distribution for the  $y$  parameter found when requiring the  $x$  to be bound between  $-5 \leq x \leq 5$  and  $0 \leq x \leq 5$  for the black and red distribution, respectively.

### 1242 4.3.2 Parameter Estimation and Credible Intervals

1243 The purpose of this analysis is to determine the best fit values for the oscillation  
1244 parameters that the beam and atmospheric samples are sensitive to:  $\sin^2(\theta_{23})$ ,  
1245  $\sin^2(\theta_{13})$ ,  $\Delta m_{32}^2$ , and  $\delta_{CP}$ . The posterior probability density, taken from the output  
1246 MCMC chain, is binned in these parameters. The parameter best-fit point is then  
1247 taken to be the value that has the highest posterior probability. This is performed  
1248 in both one and two-dimensional projections.

1249 However, the single best-fit point in a given parameter is not of much use on its  
1250 own. The uncertainty on the best-fit point must also be presented using credible  
1251 intervals. The definition of the  $1\sigma$  credible interval is that there is 68% belief  
1252 that the parameter is within those bounds. For a more generalised definition,  
1253 the credible interval is the region,  $R$ , of the posterior distribution that contains  
1254 a specific fraction of the total probability, such that

$$\int_R P(\theta|D)d\theta = \alpha. \quad (4.8)$$

1255 Where  $\theta$  is the parameter being evaluated. This technique then calculates  
1256 the  $\alpha \times 100\%$  credible interval.

1257 In practice, this analysis uses the highest posterior density (HPD) credible in-  
1258 tervals which are calculated through the following method. First, the probability  
1259 distribution is area-normalised such that it has an integrated area equal to 1.0.  
1260 The bins of probability are then summed from the highest to lowest until the sum  
1261 exceeds the  $1\sigma$  level (0.68 in this example). This process is repeated for a range of  
1262 credible intervals, notably the  $1\sigma$ ,  $2\sigma$  and  $3\sigma$  along with other levels where the  
1263 critical values for each level can be found in [75]. This process can be repeated  
1264 for the two-dimensional probability distributions by creating two-dimensional  
1265 contours of credible intervals rather than a one-dimensional result.

### 1266 4.3.3 Bayesian Model Comparisons

1267 Due to the matter resonance, this analysis has some sensitivity to the mass  
 1268 hierarchy of neutrino states (whether  $\Delta m_{32}^2$  is positive or negative) and the  
 1269 octant of  $\sin^2(\theta_{23})$ . The Bayesian approach utilised within this analysis gives an  
 1270 intuitive method of model comparison by determining which hypothesis is most  
 1271 favourable. Taking the ratio of Equation 4.3 for the two hypotheses of normal  
 1272 hierarchy,  $NH$ , and inverted hierarchy,  $IH$ , gives

$$\frac{P(\vec{\theta}_{NH}|D)}{P(\vec{\theta}_{IH}|D)} = \frac{P(D|\vec{\theta}_{NH})}{P(D|\vec{\theta}_{IH})} \times \frac{P(\vec{\theta}_{NH})}{P(\vec{\theta}_{IH})}. \quad (4.9)$$

1273 The middle term defines the Bayes factor,  $B(NH/IH)$ , which is a data-driven  
 1274 interpretation of how strong the data prefers one hierarchy to the other. For this  
 1275 analysis, equal priors on both mass hierarchy hypotheses are chosen ( $P(\vec{\theta}_{NH}) =$   
 1276  $P(\vec{\theta}_{IH}) = 0.5$ ). In practice, the MCMC chain proposes a value of  $|\Delta m_{32}^2|$  and  
 1277 then applies a 50% probability that the value is sign flipped. Consequently,  
 1278 the Bayes factor can be calculated from the ratio of the probability density in  
 1279 either hypothesis. This equates to counting the number of steps taken in the  
 1280 normal and inverted hierarchies and taking the ratio. The same approach can be  
 1281 taken to compare the upper octant (UO) compared to the lower octant (LO)  
 1282 hypothesis of  $\sin^2(\theta_{23})$ .

$\log_{10}(B_{AB})$	$B_{AB}$	Strength of Preference
< 0.0	< 1	No preference for hypothesis A (Supports hypothesis B)
0.0 – 0.5	1.0 – 3.16	Preference for hypothesis A is weak
0.5 – 1.0	3.16 – 10.0	Preference for hypothesis A is substantial
1.0 – 1.5	10.0 – 31.6	Preference for hypothesis A is strong
1.5 – 2.0	31.6 – 100.0	Preference for hypothesis A is very strong
> 2.0	> 100.0	Decisive preference for hypothesis A

**Table 4.1:** Jeffreys scale for strength of preference for two models  $A$  and  $B$  as a function of the calculated Bayes factor ( $B_{AB} = B(A/B)$ ) between the two models [137]. The original scale is given in terms of  $\log_{10}(B(A/B))$  but converted to linear scale for easy comparison throughout this thesis.

1283 Whilst the value of the Bayes factor should always be shown, the Jeffreys scale  
 1284 [137] (highlighted in Table 4.1) gives an indication of the strength of preference

1285 for one model compared to the other. Other interpretations of the strength of  
1286 preference of a model exist, e.g. the Kass and Raferty Scale [138].

#### 1287 4.3.4 Comparison of MCMC Output to Expectation

1288 To ensure the fit is performing well, a best-fit spectrum is produced using the  
1289 posterior probability distribution and compared with the data, allowing easy  
1290 by-eye comparisons to be made. A simple method of doing this is to perform a  
1291 comparison in the fitting parameters (For instance, the reconstructed neutrino  
1292 energy and lepton direction for T2K far detector beam samples) of the spectra  
1293 generated by the MCMC chain to ‘data’. This ‘data’ could be true data or some  
1294 variation of Monte Carlo prediction. This allows easy comparison of the MCMC  
1295 probability distribution to the data. To perform this,  $N$  steps from the post-burnin  
1296 MCMC chain are randomly selected. From these, the Monte Carlo prediction  
1297 at each step is generated by reweighting the model parameters to the values  
1298 specified at that step. Due to the probability density being directly correlated  
1299 with the density of steps in a certain region, parameter values close to the best  
1300 fit value are most likely to be selected.

1301 In practice, for each bin of the fitting parameters has a probability distribution  
1302 of event rates, with one entry per sampled MCMC step. This distribution is  
1303 binned where the bin with the highest probability is selected as the mean and an  
1304 error on the width of this probability distribution is calculated using the approach  
1305 highlighted in subsection 4.3.2. Consequently, the best fit distribution in the fit  
1306 parameter is not necessarily that which would be attained by reweighting the  
1307 Monte Carlo prediction to the most probable parameter values.

1308 A similar study can be performed to illustrate the freedom of the model  
1309 parameter space prior to the fit. This can be done by throwing parameter values  
1310 from the prior uncertainty of each parameter.

# 5

1311

1312

1313

## Simulation, Reconstruction, and Event Reduction

1314 As a crucial part of the oscillation analysis, an accurate prediction of the expected  
1315 neutrino spectrum at the far detector is required. This includes modeling the  
1316 flux generation, neutrino interactions, and detector effects. All of the simulation  
1317 packages required to do this are briefly described in section 5.1. The reconstruc-  
1318 tion of neutrino events in the far detector, including the `fitQun` algorithm, is  
1319 documented in section 5.2. This also includes data quality checks of the SK-V  
1320 data which the author performed for the T2K oscillation analysis presented at the  
1321 Neutrino 2020 conference [74]. Finally, section 5.3 describes the steps taken in the  
1322 SK detector to trigger on events of interest whilst removing the comparatively  
1323 large rate of cosmic ray muon events.

1324

### 5.1 Simulation

1325 In order to generate a Monte Carlo prediction of the expected event rate at  
1326 the far detector, all the processes in the beam and atmospheric fluxes, neutrino  
1327 interaction, and detector need to be modeled.

### 1328 5.1.1 Neutrino Flux

1329 The beamline simulation consists of three distinct parts: the initial hadron  
1330 interaction modeled by FLUKA [139], the target station geometry and particle  
1331 tracking performed by JNUBEAM, [140, 141] and any hadronic re-interactions  
1332 simulated by GCALOR [142]. The primary hadronic interactions are  $O(10)\text{GeV}$ ,  
1333 where FLUKA matches external cross-section data better than GCALOR [143].  
1334 However, FLUKA is not very adaptable so a small simulation is built to model  
1335 the interactions in the target and the output is then passed to JNUBEAM and  
1336 GCALOR for propagation. The hadronic interactions are tuned to data from the  
1337 NA61/SHINE [144–146] and HARP [147] experiments. The tuning is done by  
1338 reweighting the FLUKA and GCALOR predictions to match the external data  
1339 multiplicity and cross-section measurements, based on final state particle kine-  
1340 matics [143]. The culmination of this simulation package generates the predicted  
1341 flux for neutrino and antineutrino beam modes which are illustrated in Figure 3.7.

1342 The atmospheric neutrino flux is simulated by the HKKM model [50, 52]. The  
1343 primary cosmic ray flux is tuned to AMS [148] and BESS [149] data assuming  
1344 the US-standard atmosphere '76 [150] density profile and includes geomagnetic  
1345 field effects. The primary cosmic rays interact to generate pions and muons.  
1346 The interaction of these secondary particles to generate neutrinos is handled by  
1347 DPMJET-III [151] for energies above 32GeV and JAM [52, 152] for energies below  
1348 that value **DB: Question for Giles: Why different generators for above/below**  
1349 **32GeV?**. These hadronic interactions are tuned to BESS and L3 data [153, 154]  
1350 using the same methodology as the tuning of the beamline simulation. The  
1351 energy and cosine zenith predictions of  $\nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$  flux are given in Figure 2.3  
1352 and Figure 2.5, respectively. The flux is approximately symmetrical and peaked  
1353 around the horizon ( $\cos(\theta_Z) = 0.0$ ). This is because horizontally-going pions  
1354 and kaons can travel further than their vertically-going counterparts resulting  
1355 in a larger probability of decaying to neutrinos. The symmetry is broken in  
1356 lower-energy neutrinos due to geomagnetic effects, which modify the track of the  
1357 primary cosmic rays. Updates to the HKKM model are currently ongoing [155].

### 1358 5.1.2 Neutrino Interaction

1359 Once a flux prediction has been made for all three detectors, NEUT 5.4.0 [156, 157]  
1360 models the interactions of the neutrinos in the detectors. For the purposes of this  
1361 analysis, quasi-elastic (QE), meson exchange (MEC), single meson production  
1362 (PROD), coherent pion production (COH), and deep inelastic scattering (DIS)  
1363 interactions are simulated. These interaction categories can be further broken  
1364 down by whether they were propagated via a  $W^\pm$  boson in Charged Current  
1365 (CC) interactions or via a  $Z^0$  boson in Neutral Current (NC) interactions. CC  
1366 interactions have a charged lepton in the final state, which can be flavour-tagged  
1367 in reconstruction to determine the flavour of the neutrino. In contrast, NC  
1368 interactions have a neutrino in the final state so no flavour information can be  
1369 determined from the observables left in the detector after an interaction. This  
1370 is the reason why neutrinos that interact through NC modes are assumed to  
1371 not oscillate within this analysis. Both CC and NC interactions are modeled  
1372 for all the above interaction categories, other than MEC interactions which are  
1373 only modeled for CC events.

1374 As illustrated in Figure 5.1, CCQE interactions dominate the cross-section of  
1375 neutrino interactions around  $E_\nu \sim 0.5\text{GeV}$ . The NEUT implementation adopts  
1376 the Llewellyn Smith [158] model for neutrino-nucleus interactions, where the  
1377 nuclear ground state of any bound nucleons (neutrino-oxygen interactions) is  
1378 approximated by a spectral-function [159] model that simulates the effects of  
1379 Fermi momentum and Pauli blocking. The cross-section of QE interactions is  
1380 controlled by vector and axial-vector form factors parameterised by the BBBA05  
1381 [160] model and a dipole form factor with  $M_A^{QE} = 1.21\text{GeV}$  fit to external data  
1382 [161], respectively. NEUT implements the Valencia [162] model to simulate MEC  
1383 events, where two nucleons and two holes in the nuclear target are produced  
1384 (often called 2p2h interactions).

1385 For neutrinos of energy  $O(1)\text{GeV}$ , PROD interactions become dominant.  
1386 These predominantly produce charged and neutral pions although  $\gamma$ , kaon,

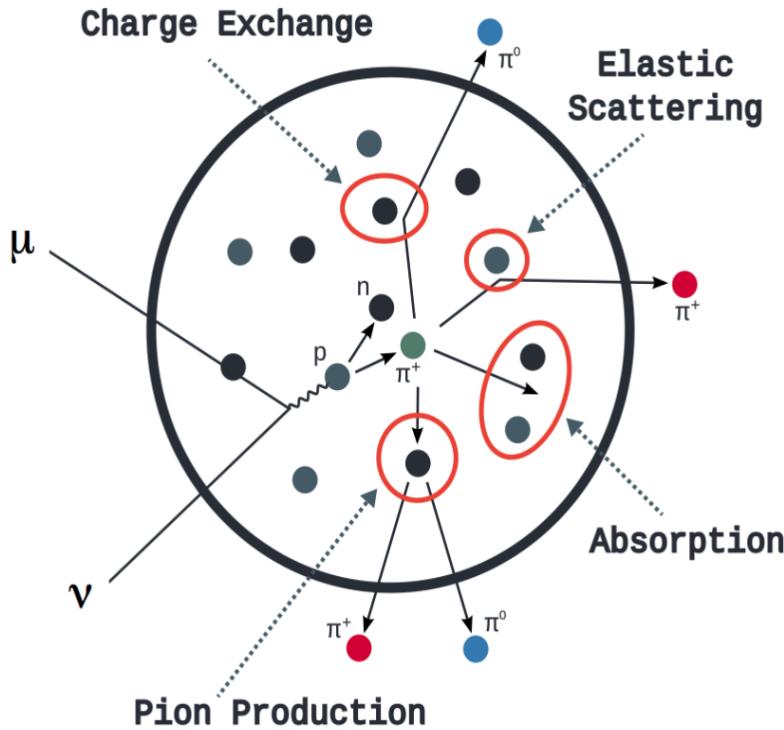


**Figure 5.1:** The NEUT prediction of the  $\nu_\mu$ -H<sub>2</sub>O cross-section overlaid on the T2K  $\nu_\mu$  flux. The charged current (black, solid) and neutral current (black, dashed) inclusive, charged current quasi-elastic (blue, solid), charged current 2p2h (blue, dashed), charged current single pion production (pink), and charged current multi- $\pi$  and DIS (Purple) cross-sections are illustrated. Figure taken from [156].

and  $\eta$  production is also considered. To simulate these interactions, the Berger-Sehgal [163] model is implemented within NEUT. It simulates the excitation of a nucleon from a neutrino interaction, production of an intermediate baryon, and the subsequent decay to a single meson or  $\gamma$ . Pions can also be produced through COH interactions, which occur when the incoming neutrino interacts with the entire oxygen nucleus leaving a single pion outside of the nucleus. NEUT utilises the Berger-Sehgal [164] model to simulate these COH interactions.

DIS and multi- $\pi$  producing interactions become the most dominant for energies  $> O(5)\text{GeV}$ . PYTHIA [165] is used to simulate any interaction with invariant mass  $W > 2\text{GeV}/c^2$ , which produces at least one meson. For any interaction which produces at least two mesons but has  $W < 2\text{GeV}/c^2$ , the

<sup>1398</sup> Bronner model is used [166]. Both of these models use Parton distribution  
<sup>1399</sup> functions based on the Bodek-Yang model [167–169].



**Figure 5.2:** Illustration of the various processes which a pion can undergo before exiting the nucleus. Taken from [170].

<sup>1400</sup> Any pion that is produced within the nucleus can re-interact through final  
<sup>1401</sup> state interactions before it exits, as illustrated by the scattering, absorption,  
<sup>1402</sup> production, and exchange interactions in Figure 5.2. These re-interactions alter  
<sup>1403</sup> the observable particles within the detector. For instance, if the charged pion  
<sup>1404</sup> from a CC PROD interaction is absorbed, the observables would mimic a CC QE  
<sup>1405</sup> interaction. To simulate these effects, NEUT uses a semi-classical intranuclear  
<sup>1406</sup> cascade model [156]. This cascade functions by stepping the pion through the  
<sup>1407</sup> nucleus in fixed-length steps equivalent to  $dx = R_N/100$ , where  $R_N$  is the radius  
<sup>1408</sup> of the nucleus. At each step, the simulation allows the pion to interact through  
<sup>1409</sup> scattering, charged exchange, absorption, or production with an interaction-  
<sup>1410</sup> dependent probability calculated from a fit to external data [171]. This cascade  
<sup>1411</sup> continues until the pion is absorbed or exits the nucleus.

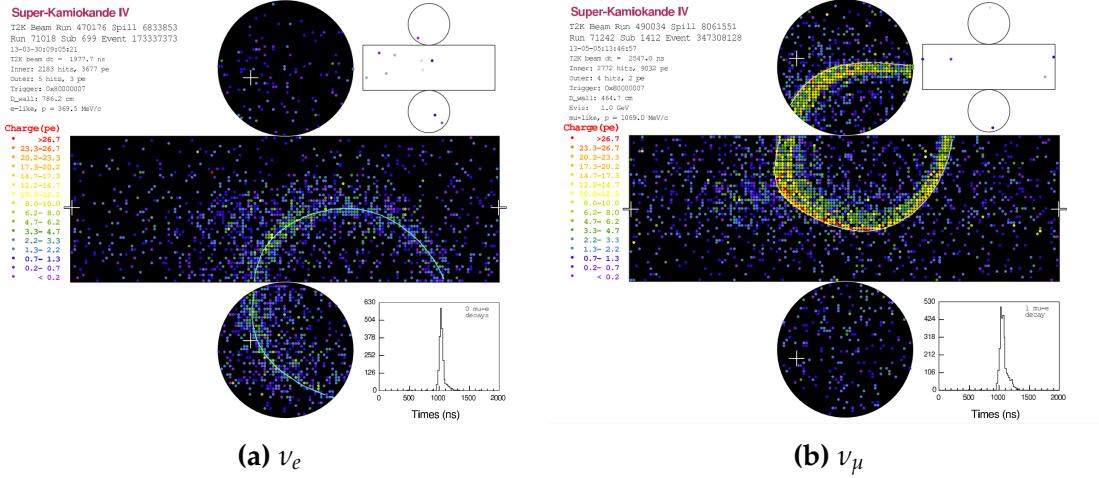
### 1412 5.1.3 Detector

1413 Once the final state particle kinematics have been determined by NEUT, they  
1414 are passed into the detector simulation. The near detectors, ND280 and INGRID,  
1415 are simulated using a GEANT4 package [115, 172] to simulate the detector geom-  
1416 etry, particle tracking, and energy deposition. The response of the detectors is  
1417 simulated using the elecSim package [115]. The far detector simulation is based  
1418 upon the original Kamiokande experiment software which uses the GEANT3-based  
1419 SKDETSIM [115, 173] package. This simulates the interactions of particles in  
1420 the water as well as Cherenkov light production. The water quality and PMT  
1421 calibration measurements detailed in subsection 3.1.2 are also used within this  
1422 simulation to make accurate predictions of the detector response.

1423 **DB: Make sure the rest of this section flows and is relevant to simulation**

1424 Any event which generates optical photons that occurs in SK will be ob-  
1425 served by the PMT array, where each PMT records the time and accumulated  
1426 charge. This recorded information is shown in event displays similar to those  
1427 illustrated in Figure 5.3. To be useful for physics analyses, this series of PMT  
1428 hit information needs to be reconstructed to determine the number and identity  
1429 of particles and their kinematics (or track parameters): four-vertex, direction,  
1430 and momentum. The reconstruction uses the fact that the charge and timing  
1431 distribution of photons generated by a particular particle in an event is dependent  
1432 upon its initial kinematics. Electron and muon rings are distinguished by their  
1433 “fuzziness”. Muons are heavier and less affected by scattering or showering  
1434 meaning they typically produce “crisp” rings. Electrons are more likely to interact  
1435 via electromagnetic showering or scattering which results in larger variations  
1436 of their direction from the initial direction. Consequently, electrons typically  
1437 produce “fuzzier” rings compared to muons.

## 1438 5.2 Event Reconstruction at SK



**Figure 5.3:** Event displays from Super Kamiokande illustrating the “crisp” ring from a muon and the typically “fuzzier” electron ring. Each pixel represents a PMT and the color scheme denotes the accumulated charge deposited on that PMT. Figures taken from [174].

1439 DB: Better separate FiTQun and APFit from the below paragraph, and add a  
 1440 little more information about APFit - Hough transform For the purposes of this  
 1441 analysis, the fiTQun reconstruction algorithm [175] is utilised. Its core function  
 1442 is to compare a prediction of the accumulated charged and timing distribution  
 1443 from each PMT, generated for a particular particle identity and track parameters,  
 1444 to that observed in the neutrino event. It determines the preferred values by  
 1445 maximising a likelihood function (or minimising a log-likelihood function) which  
 1446 includes information from PMTs which were hit and those that were not hit. The  
 1447 fiTQun algorithm is based on the key concepts of the MiniBooNE reconstruction  
 1448 algorithm [176]. It improves upon the previous APFit reconstruction algorithm  
 1449 which has been used for many previous SK analyses. APFit fits the vertex from  
 1450 timing information and then fits the momentum and direction of the particle  
 1451 from PMT hits within a 43 deg Cherenkov cone (assuming an ultra-relativistic  
 1452 particle). Due to the additional information being used in the fit, fiTQun improves  
 1453 both the accuracy of the fit parameters and the rejection of neutral current  $\pi^0$   
 1454 events as compared to APFit [177, 178].

1455 An event in SK can consist of prompt (or primary) and decay (or secondary)  
 1456 particles. For example, a charged current muon neutrino interaction can gen-

1457 erate two particles that have the potential of generating Cherenkov photons  
 1458 (assuming the proton is below the Cherenkov threshold): the prompt muon,  
 1459 and the secondary decay-electron from the muon, approximately  $2\mu\text{s}$  later. To  
 1460 reconstruct all particles within an event, it is divided into time clusters which are  
 1461 called “subevents”. Subevents after the primary subevent are considered to  
 1462 be decay electrons.

1463 The main steps of the `fitQun` reconstruction algorithm are:

- 1464 • **Vertex pre-fitting:** An estimate of the vertex is made using a goodness-of-fit  
 1465 metric based on PMT hit times
- 1466 • **Peak finding:** The initial time of each subevent is determined by clustering  
 1467 events by time residuals
- 1468 • **Single-ring fits:** Given the pre-fit vertex and estimated time of interaction,  
 1469 a maximum likelihood technique searches for a single particle generating  
 1470 light. Electron, muon, charged pion, and proton hypotheses are considered
- 1471 • **Multi-ring fits:** Seeded from the single-ring fits, hypotheses with multiple  
 1472 light-producing particles are considered using the same maximum likeli-  
 1473 hood technique. Electron-like or charged pion-like rings are added until  
 1474 the likelihood stops improving

1475 To find all the subevents in an event, a vertex goodness metric is calculated  
 1476 for some vertex position  $\vec{x}$  and time  $t$ ,

$$G(\vec{x}, t) = \sum_i^{\text{hit PMTs}} \exp \left( -\frac{1}{2} \left( \frac{T_{Res}^i(\vec{x}, t)}{\sigma} \right)^2 \right), \quad (5.1)$$

1477 where

$$T_{Res}^i(\vec{x}, t) = t^i - t - |R_{PMT}^i - \vec{x}| / c_n, \quad (5.2)$$

1478 is the residual hit time. It is the difference in time between the PMT hit time  
 1479  $t^i$ , of the  $i^{th}$  PMT, and the expected time of the PMT hit if the photon was at

1480 the vertex.  $R_{PMT}^i$  is the position of the  $i^{th}$  PMT,  $c_n$  is the speed of light in water  
 1481 and  $\sigma = 4\text{ns}$  which is comparable to the time resolution of the PMT. When the  
 1482 proposed fit values of time and vertex are close to the true values,  $T_{Res}^i(\vec{x}, t)$  tends  
 1483 to zero resulting in subevents appearing as spikes in the goodness metric. The  
 1484 proposed fit vertex and time are grid-scanned, and the values which maximise  
 1485 the goodness metric are selected as the “pre-fit vertex”. Whilst this predicts a  
 1486 vertex for use in the clustering algorithm, the final vertex is fit using the higher-  
 1487 precision maximum likelihood method described below.

1488 Once the pre-fit vertex has been determined, the goodness metric is scanned as  
 1489 a function of  $t$  to determine the number of subevents. A peak-finding algorithm  
 1490 is then used on the goodness metric, requiring the goodness metric to exceed  
 1491 some threshold and drop below a reduced threshold before any subsequent  
 1492 additional peaks are considered. The thresholds are set such that the rate of  
 1493 false peak finding is minimised while still attaining good data to Monte Carlo  
 1494 agreement. To improve performance, the pre-fit vertex for each delayed subevent  
 1495 is re-calculated after PMT hits from the previous subevent are masked. This  
 1496 improves the decay-electron tagging performance. Once all subevents have  
 1497 been determined, the time window around each subevent is then defined by the  
 1498 earliest and latest time which satisfies  $-180 < T_{Res}^i < 800\text{ns}$ . The subevents and  
 1499 associated time windows are then used as seeds for further reconstruction.

1500 For a given subevent, the `fitQun` algorithm constructs a likelihood based on  
 1501 the accumulated charge  $q_i$  and time information  $t_i$  from the  $i^{th}$  PMT,

$$L(\Gamma, \vec{\theta}) = \prod_j^{\text{unhit}} P_j(\text{unhit}|\Gamma, \vec{\theta}) \prod_i^{\text{hit}} \{1 - P_i(\text{unhit}|\Gamma, \vec{\theta})\} f_q(q_i|\Gamma, \vec{\theta}) f_t(t_i|\Gamma, \vec{\theta}). \quad (5.3)$$

1502 Where  $\vec{\theta}$  defines the track parameters; vertex position, direction vector and  
 1503 momenta, and  $\Gamma$  represents the particle hypothesis.  $P_i(\text{unhit}|\Gamma, \vec{\theta})$  is the proba-  
 1504 bility of the  $i^{th}$  tube to not register a hit given the track parameters and particle  
 1505 hypothesis. The charge likelihood,  $f_q(q_i|\Gamma, \vec{\theta})$ , and time likelihood,  $f_t(t_i|\Gamma, \vec{\theta})$ ,

represents the probability density function of observing charge  $q_i$  and time  $t_i$  on the  $i^{th}$  PMT given the specified track parameters and particle hypothesis.

The predicted charge is calculated based on contributions from both the direct light and the scattered light. The direct light contribution is determined based on the integration of the Cherenkov photon profile along the track. PMT angular acceptance, water quality, and calibration measurements discussed in subsection 3.1.2 are included to accurately predict the charge probability density at each PMT. The scattered and reflected light is calculated in a similar way, although it includes a scattering function that depends on the vertex of the particle and the position of the PMT. The charge likelihood is calculated by comparing the prediction to the observed charge in the PMT which is tuned to the PMT simulation.

The time likelihood is approximated to depend on the vertex  $\vec{x}$ , direction  $\vec{d}$ , and time  $t$  of the track as well as the particle hypothesis. The expected time for PMT hits is calculated by assuming unscattered photons being emitted from the midpoint of the track,  $S_{mid}$ ,

$$t_{exp}^i = t + S_{mid}/c + |R_{PMT}^i - \vec{x} - S_{mid}\vec{d}|/c_n, \quad (5.4)$$

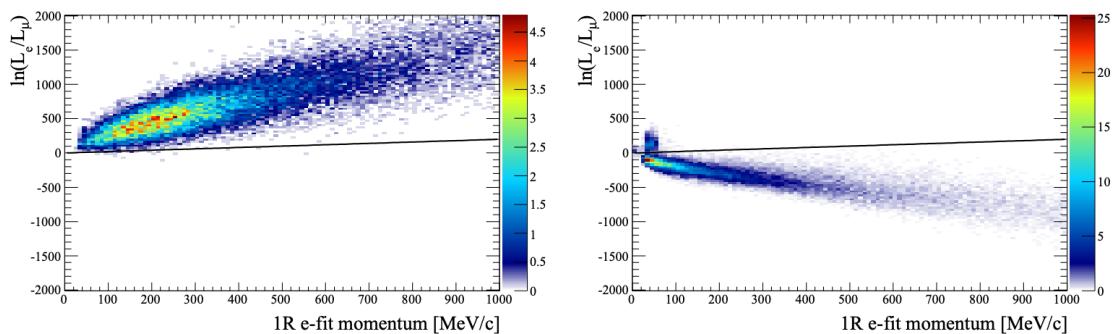
where  $c$  is the speed of light in a vacuum. The time likelihood is then expressed in terms of the residual difference between the PMT hit time and the expected hit time,  $t_{Res}^i = t^i - t_{exp}^i$ . The particle hypothesis and momentum also affect the Cherenkov photon distribution. These parameters modify the shape of the time likelihood density since in reality not all photons are emitted at the midpoint of the track. As with the charge likelihood, the contributions from both the direct and scattered light to the time likelihood density are calculated separately, which are both calculated from particle gun Monte Carlo studies.

The track parameters and particle identity which maximise  $L(\Gamma, \vec{\theta})$  are defined as the best-fit parameters. In practice MINUIT [179] is used to minimise the value of  $-\ln L(\Gamma, \vec{\theta})$ . The `fitQun` algorithm considers an electron-like, muon-like, and charged pion-like hypothesis for events with a single final state particle, denoted

<sup>1534</sup> “single-ring events”. The particle’s identity is determined by taking the ratio of  
<sup>1535</sup> the likelihood of each of the hypotheses. For instance, electrons and muons are  
<sup>1536</sup> distinguished by considering the value of  $\ln(L(e, \vec{\theta}_e)/L(\mu, \vec{\theta}_\mu))$  in comparison  
<sup>1537</sup> to the reconstructed momentum of the electron hypothesis, as illustrated by  
<sup>1538</sup> Figure 5.4. The coefficients of the discriminator between electron-like and muon-  
<sup>1539</sup> like events are determined from Monte Carlo studies [175]. Similar distributions  
<sup>1540</sup> exist for distinguishing electron-like events from  $\pi^0$ -like events, and muon-like  
<sup>1541</sup> events from pion-like events. The cuts are defined as,

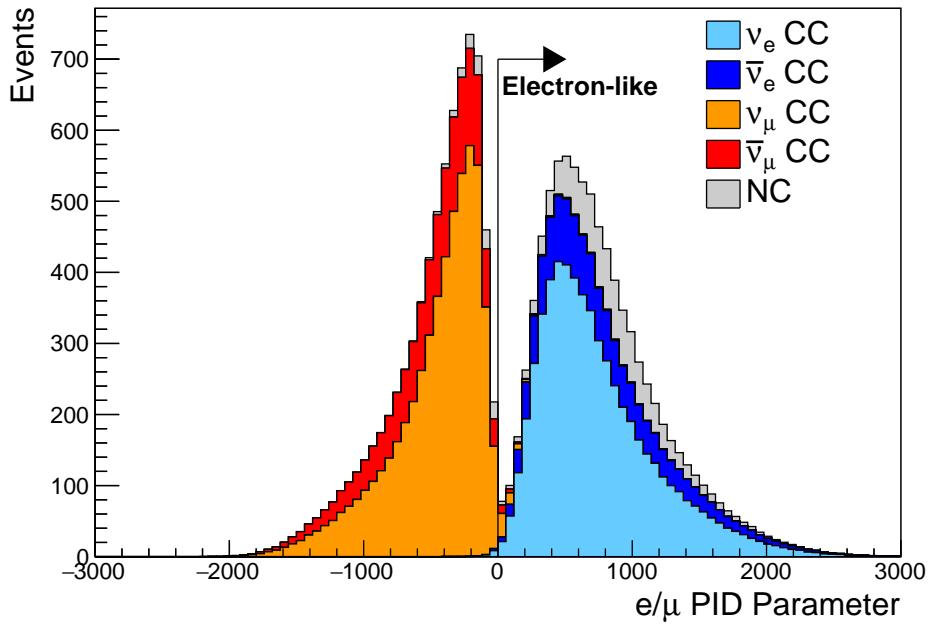
$$\begin{aligned} \text{Electron/Muon} : & \ln(L_e/L_\mu) > 0.2 \times p_e^{rec} [\text{MeV}], \\ \text{Electron}/\pi^0 : & \ln(L_e/L_{\pi^0}) < 175 - 0.875 \times m_{\gamma\gamma} [\text{MeV}], \\ \text{Muon/Pion} : & \ln(L_\mu/L_{\pi^\pm}) < 0.15 \times p_\mu^{rec} [\text{MeV}], \end{aligned} \quad (5.5)$$

<sup>1542</sup> as taken from [180], where  $p_e^{rec}$  and  $p_\mu^{rec}$  are the reconstructed momentum of the  
<sup>1543</sup> single-ring electron and muon fits, respectively.  $m_{\gamma\gamma}$  represents the reconstructed  
<sup>1544</sup> invariant mass of the two photons emitted from  $\pi^0$  decay. Typically, the distance  
<sup>1545</sup> between a particular entry in these two-dimensional distributions and the cut-line  
<sup>1546</sup> is termed the PID parameter and is illustrated in Figure 5.5.



**Figure 5.4:** The difference of the electron-like and muon-like log-likelihood compared to the reconstructed single-ring fit momentum for atmospheric  $\nu_e$  (left) and  $\nu_\mu$  (right) samples. The black line represents the cut used to discriminate electron-like and muon-like events, with coefficients obtained from Monte Carlo studies. Figures from [175].

<sup>1547</sup> The `fitQun` algorithm also considers a  $\pi^0$  hypothesis. To do this, it performs  
<sup>1548</sup> a fit looking for two standard electron-hypothesis tracks which point to the



**Figure 5.5:** The electron/muon PID separation parameter for all sub-GeV single-ring events in SK-IV. The charged current (CC) component is broken down in four flavours of neutrino ( $\nu_\mu$ ,  $\bar{\nu}_\mu$ ,  $\nu_e$  and  $\bar{\nu}_e$ ). Events with positive values of the parameter are determined to be electron-like.

1549 same four-vertex. This assumes the electron tracks are generated from photon-  
 1550 conversion so the electron tracks actually appear offset from the proposed  $\pi^0$   
 1551 vertex. For these fits, the conversion length, direction, and momentum of each  
 1552 photon are also considered as track parameters which are then fit in the same  
 1553 methodology as the standard single-ring hypotheses.

1554 Whilst lower energy events are predominantly single-ring events, higher  
 1555 energy neutrino events can generate final states with multiple particles which  
 1556 generate Cherenkov photons. These “multi-ring” hypotheses are also considered  
 1557 in the `fitQun` algorithm. When calculating the charge likelihood density, the  
 1558 predicted charge associated with each ring is calculated separately and then  
 1559 summed to calculate the total accumulated charge on each PMT. Similarly,  
 1560 the time likelihood for the multi-ring hypothesis is calculated assuming each  
 1561 ring is independent. Each track is time-ordered based on the time of flight  
 1562 from the center of the track to the PMT and the direct light from any ring

1563 incident on the PMT is assumed to arrive before any scattered light. To reduce  
1564 computational resource usage, the multi-ring fits only consider electron-like  
1565 and charged pion-like rings as the pion fit can be used as a proxy for a muon  
1566 fit due to their similar mass.

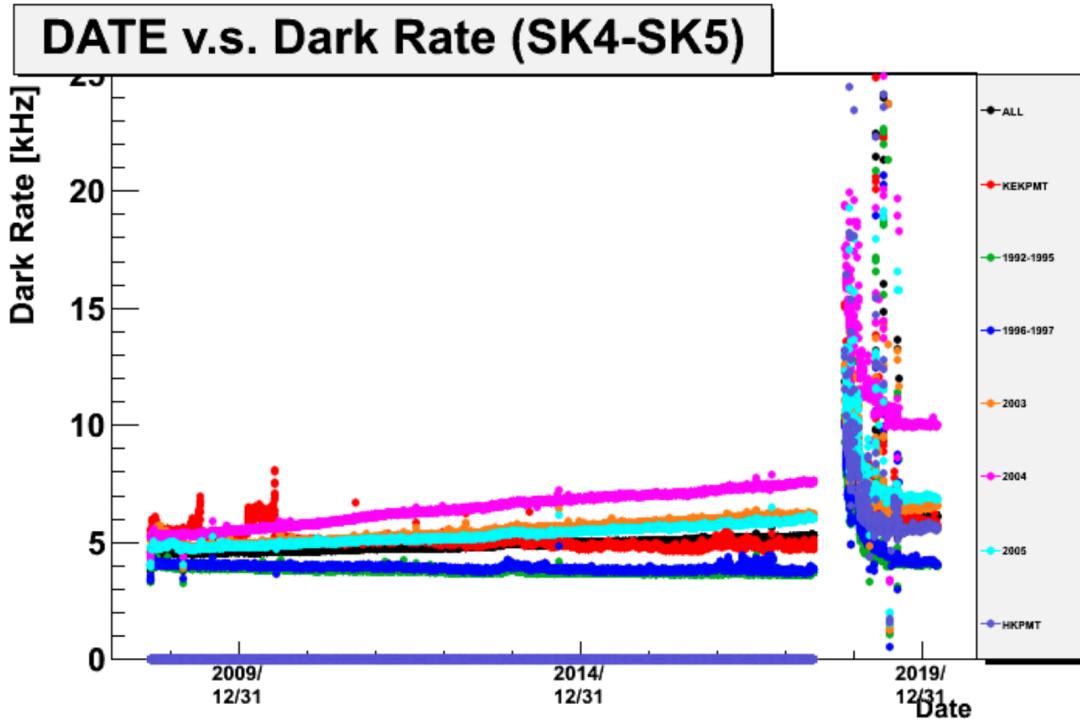
1567 **DB: Figure out how pion-like rings are different from the muon hypothesis**  
1568 Multi-ring fits proceed by proposing another ring to the previous fit and  
1569 then fitting the parameters in the method described above. Typically, multi-ring  
1570 fits have the largest likelihood because of the additional degrees of freedom  
1571 introduced. A likelihood value is calculated for the  $n$ -ring and  $(n + 1)$ -ring  
1572 hypotheses, where the additional ring is only included if the likelihood value  
1573 is above 9.35, based on Monte Carlo studies in [181].

### 1574 5.2.1 Validation of Reconstruction in SK-V

1575 **DB: Link this into the previous section**

1576 As an example of how the reconstruction depends on the detector conditions,  
1577 the author of this thesis assessed the quality of `fitQun` event reconstruction for  
1578 SK-V data. The detector systematics used in the T2K-only oscillation analysis  
1579 are determined using data-to-Monte Carlo comparisons of the SK-IV data [182].  
1580 Due to tank-open maintenance occurring between SK-IV and SK-V, the dark  
1581 rate of each PMT was observed to increase in SK-V due to light exposure for a  
1582 significant time during the repairs. This increase can be seen in Figure 5.6. Run-10  
1583 of the T2K experiment was conducted in the SK-V period, so the consistency  
1584 of SK-IV and SK-V data needs to be studied to determine whether the SK-IV-  
1585 defined systematics can be applied to the run-10 data. This comparison study was  
1586 performed using the stopping muon data set for both the SK-IV and SK-V periods.  
1587 This data sample is used due to the high rate of interactions ( $O(200)$  events per  
1588 hour) as well as having similar energies to muons from CCQE  $\nu_\mu$  interactions from  
1589 beam interactions. The rate of cosmic muons does depend on the solar activity  
1590 cycle [183] but has been neglected in this comparison study. This is because the  
1591 shape of the distributions is most important for the purposes of being compared

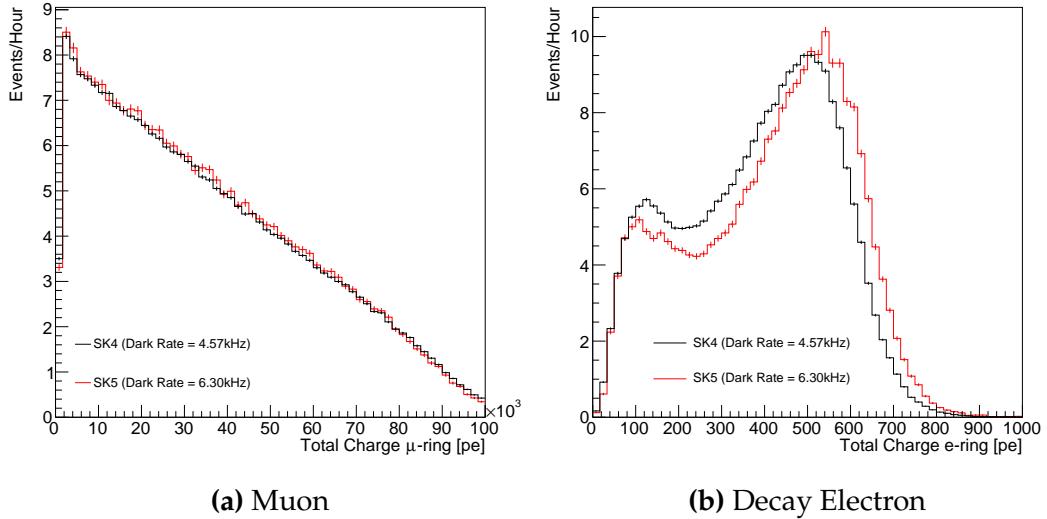
1592 to the detector systematics. The SK-IV and SK-V data samples consist of 2398.42  
1593 and 626.719 hours of data which equates to 686k and 192k events respectively.  
1594 These samples do not correspond to the full data sets of either period but do  
1595 contain enough events to be systematics limited rather than statistics limited.



**Figure 5.6:** The variation of the measured dark rate as a function of date, broken down by PMT type. The SK-IV and SK-V periods span September 2008 to May 2018 and January 2019 to July 2020, respectively. The break in measurement in 2018 corresponds to the period of tank repair and refurbishment. Figure adapted from [182].

1596 The predicted charge calculated in the `fitQun` algorithm includes a contribu-  
1597 tion from the photoelectron emission due to dark noise. Therefore, the increase  
1598 in the SK-V dark rate needs to be accounted for. In practice, the average dark rate  
1599 in each SK period is calculated and used as an input in the reconstruction. This is  
1600 calculated by averaging the dark rate per run for each period separately, using  
1601 the calibration measurements detailed in subsection 3.1.2. The average dark rate  
1602 from SK-IV and SK-V were found to be 4.57kHz and 6.30kHz, respectively. The  
1603 charges associated with the muon and decay electron subevents are illustrated in  
1604 Figure 5.7. The photoelectron emission from dark noise is more significant for

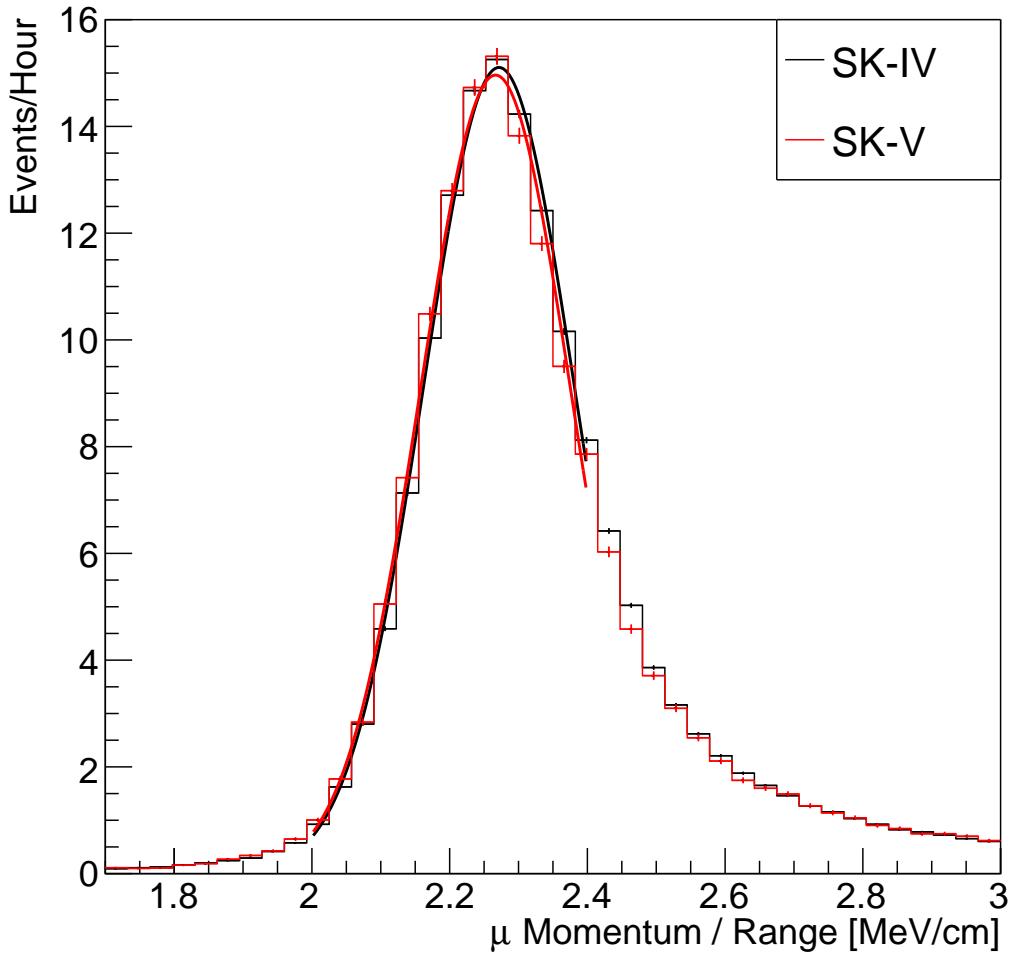
events that have lower energy. This is because this contribution becomes more comparable to the number of photoelectrons emitted from incident photons in lower-energy events. This behaviour is observed in the data, where the charge deposited by the muon subevent is mostly unaffected by the increase in dark rate, whilst the charge associated with the decay-electron is clearly affected.



**Figure 5.7:** Comparison of the measured raw charge deposited per event from the stopping muon data samples between SK-IV (Blue) and SK-V (Red), split by the primary muon subevent (left) and the associated decay electron subevent (right).

The energy scale systematic is estimated from data-to-Monte Carlo differences in the stopping muon sample in [184] and found to be 2.1%. To determine the consistency of SK-IV and SK-V with respect to the energy scale systematic, the muon momentum distribution is compared between the two SK periods. As the total number of Cherenkov photons is integrated across the track length, the reconstructed momentum divided by track length (or range) is compared between SK-IV and SK-V as illustrated in Figure 5.8.

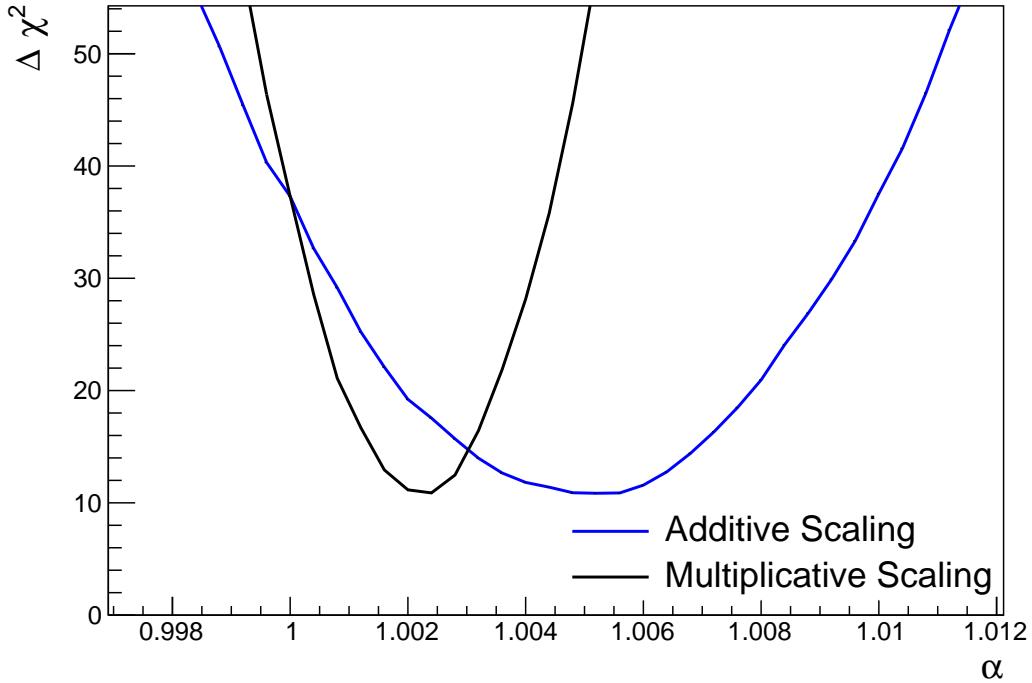
The consistency between these muon distributions has been computed in two ways. Firstly, a Gaussian is fit to the peak of each distribution separately, whose mean is found to be  $(2.272 \pm 0.003)\text{MeV/cm}$  and  $(2.267 \pm 0.006)\text{MeV/cm}$  for SK-IV and SK-V respectively. The ratio of these is equal to  $1.002 \pm 0.003$ . The means of the Gaussian fits are consistent with the expected stopping power of a minimum



**Figure 5.8:** The distribution of the reconstructed momentum from the muon ring divided by the distance between the reconstructed muon and decay electron vertices as found in the stopping muon data sets of SK-IV (Black) and SK-V (Red). Only events with one tagged decay electron are considered. A Gaussian fit is considered in the range  $[2.0, 2.4]$  MeV/cm and illustrated as the solid curve.

<sub>1622</sub> ionising muon for a target material (water) with  $Z/A \sim 0.5$  [185]. The second  
<sub>1623</sub> consistency check is performed by introducing a nuisance parameter,  $\alpha$ , which  
<sub>1624</sub> modifies the SK-V distribution. The value of  $\alpha$  which minimises the  $\chi^2$  value  
<sub>1625</sub> between the SK-IV and SK-V is determined by scanning across a range of values.  
<sub>1626</sub> This is repeated by applying the nuisance parameter as both a multiplicative  
<sub>1627</sub> factor and an additive shift. The  $\chi^2$  distributions for different values of  $\alpha$  is  
<sub>1628</sub> illustrated in Figure 5.9. The values which minimise the  $\chi^2$  are found to be 0.0052

<sup>1629</sup> and 1.0024 for the additive and multiplicative implementations, respectively. No  
<sup>1630</sup> evidence of shifts larger than the 2.1% uncertainty on the energy scale systematic  
<sup>1631</sup> has been found in the reconstructed momentum distribution of SK-IV and SK-V.

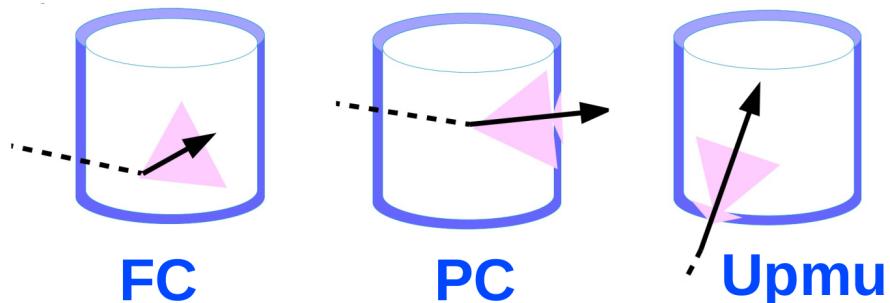


**Figure 5.9:** The  $\chi^2$  difference between the SK-IV and SK-V reconstructed muon momentum divided by range when the SK-V is modified by the scaling parameter  $\alpha$ . Both additive (Blue) and multiplicative (Black) scaling factors have been considered. In practice, the additive scaling factor actually uses the value of  $(\alpha - 1.0)$  but is illustrated like this so the results can be shown on the same axis range.

### 5.3 Event Reduction at SK

<sup>1632</sup> In normal data-taking operations, the SK detector observes many background  
<sup>1633</sup> events alongside the beam and atmospheric neutrino signal events of physics  
<sup>1634</sup> interest for this thesis. Cosmic ray muons and flasher events, which are the spon-  
<sup>1635</sup> taneous discharge of a given PMT, contribute the largest amount of background  
<sup>1636</sup> events in the energy range relevant to this thesis. Therefore the data recorded  
<sup>1637</sup> is reduced with the aim of removing these background events. The reduction  
<sup>1638</sup> process is detailed in [54, 85] and briefly summarised below.  
<sup>1639</sup>

Atmospheric neutrino events observed in the SK detector are categorised into three different types of samples: fully contained (FC), partially contained (PC) and up-going muon (Up- $\mu$ ), using PMT hit signatures in the inner and outer detector (ID and OD, respectively). To identify FC neutrino events, it is required that the neutrino interacts inside the fiducial volume of the ID and that no significant OD activity is observed. For this analysis, an event is defined to be in the fiducial volume provided the event vertex is at least 0.5m away from the ID walls. PC events have the same ID requirements but can have a larger signal present inside the OD. Typically, only high energy muons from  $\nu_\mu$  interactions can penetrate the ID wall. The Up- $\mu$  sample contains events where muons are created from neutrino interactions in the OD water or rock below the tank. They then propagate upwards through the detector. Downward-going muons generated from neutrino interactions above the tank are neglected because of the difficulty in separating their signature from the cosmic muon shower background. The sample categories are visually depicted in Figure 5.10.



**Figure 5.10:** A depiction of the topology patterns for fully-contained (FC), partially-contained (PC), and up-going muon (Up- $\mu$ ) samples included in this analysis.

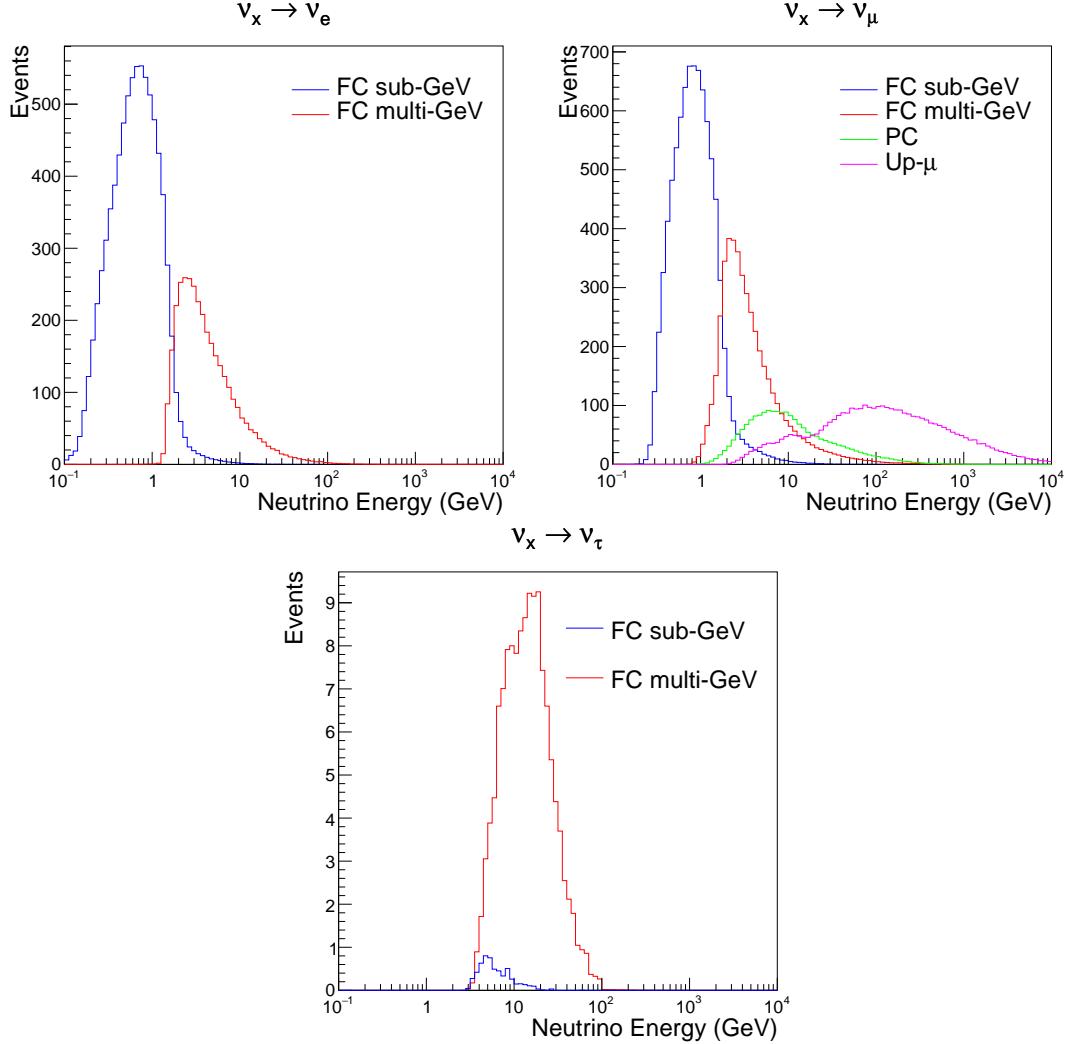
Based on the event characteristics, as defined by the `fitQun` event reconstruction software, the FC events are categorised by

- **Visible Energy:** equal to the sum of the reconstructed kinetic energy of particles above the Cerenkov threshold for all rings present in the event. The purpose is to separate events into sub-GeV and multi-GeV categories.

- **Number of observed Cerenkov rings.** The purpose is to separate single-ring and multi-ring events, where single-ring events predominantly consist of quasi-elastic interactions and multi-ring events are typically resonant pion production or deep inelastic scattering events.
- **Particle identification parameter of the most energetic ring:** A value determined from the maximum likelihood value based on `fitQun`'s electron, muon, or pion hypothesis. The purpose is to separate electron-like and muon-like events.
- **Number of decay electrons:** The purpose is to separate quasi-elastic events (which have one decay electron emitted from the muon decay) and resonant pion production events (which have two decay electrons emitted from the muon and pion).

The PC and Up- $\mu$  categories are broken down into “through-going” and “stopping” samples depending on whether the muon leaves the detector. This is because the PC stopping events deposit the entire energy of the interaction into the detector, resulting in better reconstruction. The energy of events that exit the detector has to be estimated, with a typically worse resolution, which introduces much larger systematic uncertainties. Through-going Up- $\mu$  samples are further broken down by whether any hadronic showering was observed in the event which typically indicates DIS interactions. The expected neutrino energy for the different categories is given in Figure 5.11. FC sub-GeV and multi-GeV events peak around 0.7GeV and 3GeV respectively, with slightly different peak energies for  $\nu_e$  and  $\nu_\mu$  oscillation channels. PC and Up- $\mu$  are almost entirely comprised of  $\nu_\mu$  events and peak around 7GeV and 100GeV, respectively.

The first two steps in the FC reconstruction remove the majority of cosmic ray muons by requiring a significant amount of ID activity compared to that measured in the OD. Events that pass this cut are typically very high momentum muons or events that leave very little activity in the OD. Consequently, a third reduction step is then applied to select cosmic-ray muons that pass the initial



**Figure 5.11:** The predicted neutrino flux of the fully contained (FC) sub-GeV and multi-GeV, partially contained (PC), and upward-going muon (Up- $\mu$ ) events. The prediction is broken down by the  $\nu_x \rightarrow \nu_e$  prediction (top left),  $\nu_x \rightarrow \nu_\mu$  prediction (top right) and  $\nu_x \rightarrow \nu_\tau$  prediction (bottom).  $\nu_x$  represents the flavours of neutrinos produced in the cosmic ray showers (electron and muon). Asimov A oscillation parameters are assumed (given in Table 2.2).

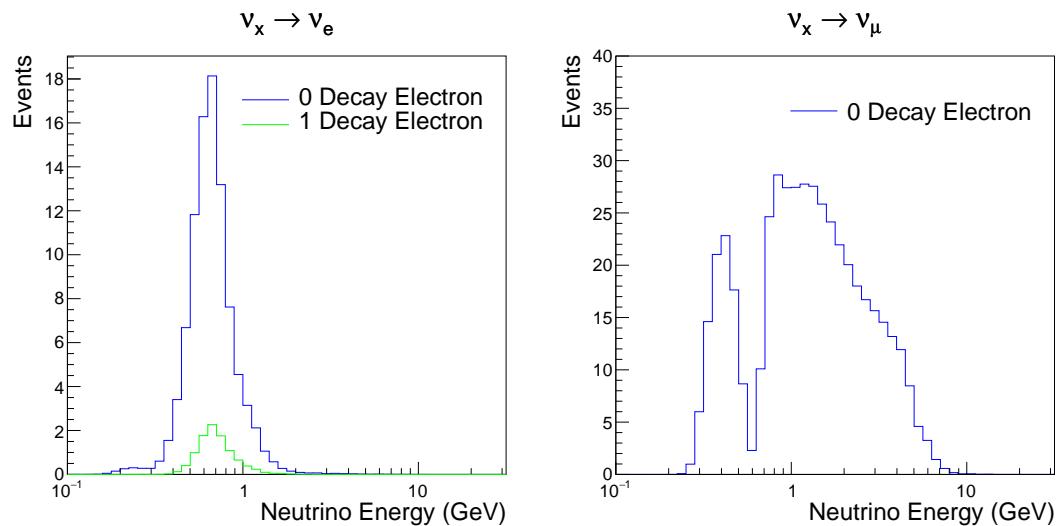
reduction step. A purpose-built cosmic muon fitter is used to determine the entrance (or exit) position of the muon and a cut is applied to OD activity contained within 8m of this position. Flasher events are removed in the fourth reduction step which is based on the close proximity of PMT hits surrounding the PMT producing the flash. Events that pass all these reduction steps are reconstructed with the APFit algorithm. The fifth step of the reduction uses information from the more precise fitter to repeat the previous two steps with

1696 tighter cuts. Muons below the Cherenkov threshold can not generate optical  
1697 photons in the ID but the associated decay electron can due to its lower mass.  
1698 These are the types of events targeted in the fifth reduction step. The final  
1699 cuts require the event vertex to be within the fiducial volume (0.5m from the  
1700 wall although the nominal distance is 2.0m), visible energy  $E_{vis} > 30\text{MeV}$  and  
1701 fewer than 16 hits within the higher energy OD cluster. The culmination of  
1702 the fully contained reduction results in 8.09 events/day in the nominal fiducial  
1703 volume [81]. The uncertainty in the reconstruction is calculated by comparing  
1704 Monte Carlo prediction to data. The largest discrepancy is found to be 1.3%  
1705 in the fourth reduction step.

1706 The PC and Up- $\mu$  events are processed through their own reduction processes  
1707 detailed in [54]. Both of these samples are reconstructed with the APFit algorithm  
1708 rather than fiTQun. This is because the efficiency of reconstructing events that  
1709 leave the detector has not been sufficiently studied for reliable systematic uncer-  
1710 tainties with fiTQun. The PC and Up- $\mu$  samples acquire events at approximately  
1711 0.66 and 1.44 events/day.

1712 Beam neutrinos events undergo the same reduction steps as FC events and are  
1713 then subject to further cuts [186]. The GPS system that links the timing between  
1714 the beam facility and SK needs to be operating correctly and there should be no  
1715 activity within the detector in the previous  $100\mu\text{s}$  before the trigger. The events  
1716 then need to triggered between  $-2\mu\text{s}$  and  $10\mu\text{s}$  of the expected spill timing.

1717 The beam neutrino samples are not split by visible energy since their energy  
1718 range is smaller than the atmospheric neutrino events. Following the T2K  
1719 analysis in [74], only single-ring beam neutrino events are considered. Similar to  
1720 atmospheric event selection, the number of decay electrons is used as a proxy for  
1721 distinguishing CCQE and CCRES events. The expected neutrino energy, broken  
1722 down by the number of decay electrons, is given in Figure 5.12.



**Figure 5.12:** The predicted flux of beam neutrinos, as a function of neutrino energy. The predictions are broken down by the number of decay electrons associated with the particular events. Asimov A oscillation parameters are assumed (given in Table 2.2).

# 6

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1724

## Sample Selections and Systematics

1725 The oscillation analysis presented within this thesis is built upon a simultaneous  
1726 fit to atmospheric data at SK, neutrino beam data in the near detector, and  
1727 beam data measured at SK. The definitions of these samples are documented  
1728 in section 6.1, section 6.2, and section 6.3, respectively. The data collected and  
1729 used within this analysis is detailed in Table 6.1. The near and far detector data  
1730 corresponds to T2K runs 2-9 and runs 1-10, respectively. The accumulated POT  
1731 and beam power for runs 1 – 10 are illustrated in Figure 6.1.

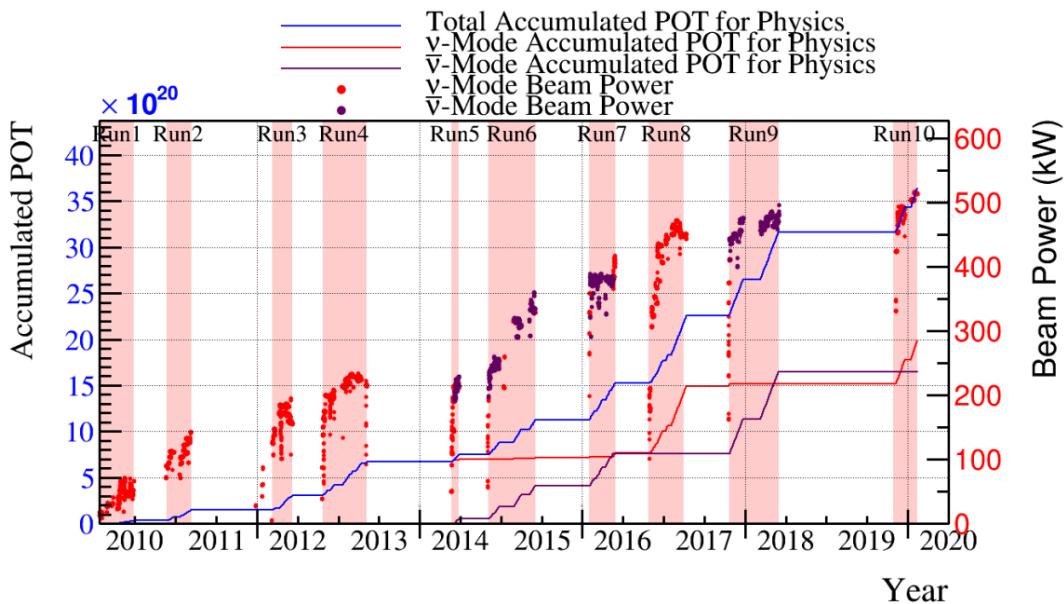
Data Type	Total
Near Detector FHC	$1.15 \times 10^{21}$ POT
Near Detector RHC	$8.34 \times 10^{20}$ POT
Far Detector FHC	$1.97 \times 10^{21}$ POT
Far Detector RHC	$1.63 \times 10^{21}$ POT
Atmospheric SK-IV	3244.4 days

**Table 6.1:** The amount of data collected in each detector used within this analysis. The data collected at the near and far detector, for both neutrino beam (FHC) and antineutrino beam (RHC), is measured as the number of protons on target (POT).

1732 The difference in POT recorded at the near and far detector is due to the  
1733 difference in downtime. The SK detector is very stable with almost 100% of  
1734 data recorded during beam operation. Due to various technical and operational

<sup>1735</sup> issues, the downtime of the near detector is significantly higher due to its more  
<sup>1736</sup> complex design and operating requirements.

<sup>1737</sup> The systematic parameters invoked within the flux, detector, and interaction  
<sup>1738</sup> models used within this analysis are documented in section 6.4. The standard  
<sup>1739</sup> configuration of the joint beam and atmospheric data fit utilises far detector sys-  
<sup>1740</sup> tematics provided in the official inputs from the two experiments. Additionally,  
<sup>1741</sup> a correlated detector model which fits the parameters used in sample selections  
<sup>1742</sup> to data has been developed and documented in subsection 6.4.5.



**Figure 6.1:** The accumulated beam data, measured as the number of protons on target (POT). The total data (blue) is given which comprises of the neutrino beam (red) and antineutrino (purple) components. The beam power for neutrino and antineutrino beams is given as the markers using the same colour scheme. The timescale runs from Run 1 which started in January 2010 until Run 10 which ended in February 2020. The ratio of accumulated data in neutrino and antineutrino beam is 54.7% : 45.3%.

<sup>1743</sup> DB: Say that I have merged T2K and SK which hadn't previously been done

## 1744 6.1 Atmospheric Samples

1745 The atmospheric event selection follows the official SK-IV analysis presented  
1746 in [85] and is documented below. The Monte Carlo prediction used within this  
1747 analysis corresponds to 500 years worth of neutrino events, which is scaled down  
1748 to match the SK-IV livetime of 3244.4 days.

1749 The fully contained (FC), partially contained (PC), and upward going muon  
1750 events ( $\text{up-}\mu$ ) which pass the reduction cuts discussed in section 5.3 are further  
1751 broken down into different samples based on reconstruction information. This  
1752 section details the samples used within this oscillation analysis, alongside the  
1753 chosen binning.

1754 FC events are first separated by the visible energy deposited within the  
1755 detector. This is calculated as the sum of the reconstructed kinetic energy  
1756 above the Cherenkov threshold for all rings present in the event. Events are  
1757 separated by whether they were above or below  $E_{\text{vis}} = 1.33\text{GeV}$ . This separates  
1758 “subGeV” and “multiGeV” events. Typically, lower energy events consist of  
1759 charged current quasi-elastic (CCQE) interactions which are better understood  
1760 and simpler to reconstruct resulting in smaller systematic uncertainties. Events  
1761 are further separated by the number of rings associated with the event due to  
1762 similar reasoning. As the oscillation probability is dependant upon the flavour  
1763 of neutrino, electron and muon events are separated using a similar likelihood  
1764 method to that discussed in section 5.2. To reduce computational resources  
1765 required for the reconstruction, only electron and pion hypotheses are considered  
1766 so this separation cut depends on the ratio of the electron to pion likelihoods,  
1767  $\log(L_e/L_\pi)$ . Finally, the number of decay electrons is used to classify events.  
1768 Charged current resonant pion production (CCRES) interactions generate a final-  
1769 state pion. This can decay, mostly likely through a muon, into a decay electron.  
1770 Therefore any electron-like event with one decay electron or muon-like event  
1771 with two decay electrons was most likely produced by a CCRES interaction.  
1772 Consequently, the number of decay electrons can be used to distinguish CCQE

<sub>1773</sub> and CCRES interaction modes. Ultimately, FC subGeV events are separated  
<sub>1774</sub> into the samples listed in Table 6.2.

Sample Name	Description
SubGeV- <i>e</i> like-0dcy	Single ring <i>e</i> -like events with zero decay electrons
SubGeV- <i>e</i> like-1dcy	Single ring <i>e</i> -like events with one or more decay electrons
SubGeV- <i>μ</i> like-0dcy	Single ring <i>μ</i> -like events with zero decay electrons
SubGeV- <i>μ</i> like-1dcy	Single ring <i>μ</i> -like events with one decay electrons
SubGeV- <i>μ</i> like-2dcy	Single ring <i>μ</i> -like events with two or more decay electrons
SubGeV- <i>π</i> 0like	Two <i>e</i> -like ring events with zero decay electrons and reconstructed $\pi^0$ mass $85 \leq m_{\pi^0} < 215$ MeV

**Table 6.2:** The fully contained subGeV samples, defined as events with visible energy  $E_{vis} < 1.33$  GeV, used within this oscillation analysis.

<sub>1775</sub> In addition to the cuts discussed above, multiGeV samples also have addi-  
<sub>1776</sub> tional cuts to separate samples which target neutrino and antineutrino events.  
<sub>1777</sub> As discussed in section 7.1, the matter resonance only occurs for neutrinos in the  
<sub>1778</sub> normal hierarchy and antineutrinos in the inverted mass hierarchy. Therefore,  
<sub>1779</sub> having flavour-enriched samples aids in the determination of the mass hierarchy.  
<sub>1780</sub> For a CCRES interaction,

$$\begin{aligned}
 \bar{\nu}_e + N &\rightarrow e^+ + N' + \pi^-, \\
 \nu_e + N &\rightarrow e^- + N' + \pi^+ \\
 &\quad \downarrow \mu^+ + \nu_\mu \\
 &\quad \downarrow e^+ + \nu_e + \bar{\nu}_\mu.
 \end{aligned} \tag{6.1}$$

<sub>1781</sub> The  $\pi^-$  emitted from a  $\bar{\nu}_e$  interaction is more likely to be captured by an  
<sub>1782</sub> oxygen nucleus than the  $\pi^+$  from  $\nu_e$  interactions [187]. These pions then decay,  
<sub>1783</sub> mostly through muons, to electrons. Therefore the number of tagged decay  
<sub>1784</sub> electrons associated with an event gives an indication of whether the interaction  
<sub>1785</sub> was due to a neutrino or antineutrino: zero for  $\bar{\nu}_e$  events, and one for  $\nu_e$  events.  
<sub>1786</sub> The ability to separate neutrino from antineutrino events is illustrated in Table 6.4,  
<sub>1787</sub> where the MultiGeV-*e*like-nue has 78% purity of CC neutrino interactions with  
<sub>1788</sub> only 7% antineutrino background, the rest consisting of NC backgrounds.

1789     The number of decay electrons discriminator works reasonably well for single-  
 1790     ring events. However, this is not the case for multi-ring events. A multiGeV  
 1791     multiring electron-like (MME) likelihood cut was introduced in [188, 189]. This  
 1792     is a two-stage likelihood selection cut. Four observables are used in the first  
 1793     likelihood cut to distinguish  $CC\nu_e$  and  $CC\bar{\nu}_e$  events from background:

- 1794       • The number of decay electrons  
 1795       • The maximum distance between the vertex of the neutrino and the decay  
 1796        electrons  
 1797       • The energy deposited by the highest energy ring  
 1798       • The particle identification of that highest energy ring

1799     Background events consist of  $CC\nu_\mu$  and NC interactions. Typically, the  
 1800     majority of the energy in these background events is carried by the hadronic  
 1801     system. Additionally, muons tend to travel further than the pions from  $CC\nu_e$   
 1802     before decaying. Thus, the parameters used within the likelihood cut target these  
 1803     typical background interaction kinematics.

Sample Name	Description
MultiGeV-elike-nue	Single ring $e$ -like events with zero decay electrons
MultiGeV-elike-nuebar	Single ring $e$ -like events with one or more decay electrons
MultiGeV-mulike	Single ring $\mu$ -like events
MultiRing-elike-nue	Two or more ring events with leading energy $e$ -like ring and passed both MME and $\nu/\bar{\nu}$ separation cuts
MultiRing-elike-nuebar	Two or more ring events with leading energy $e$ -like ring and passed MME and failed $\nu/\bar{\nu}$ separation cuts
MultiRing-mulike	Two or more ring events with leading energy $\mu$ -like ring and only requires $E_{vis} > 0.6\text{GeV}$ <b>DB: Why is this not }1.33\text{GeV}</b>
MultiRing-Other1	Two or more ring events with leading energy $e$ -like ring and failed the MME likelihood cut

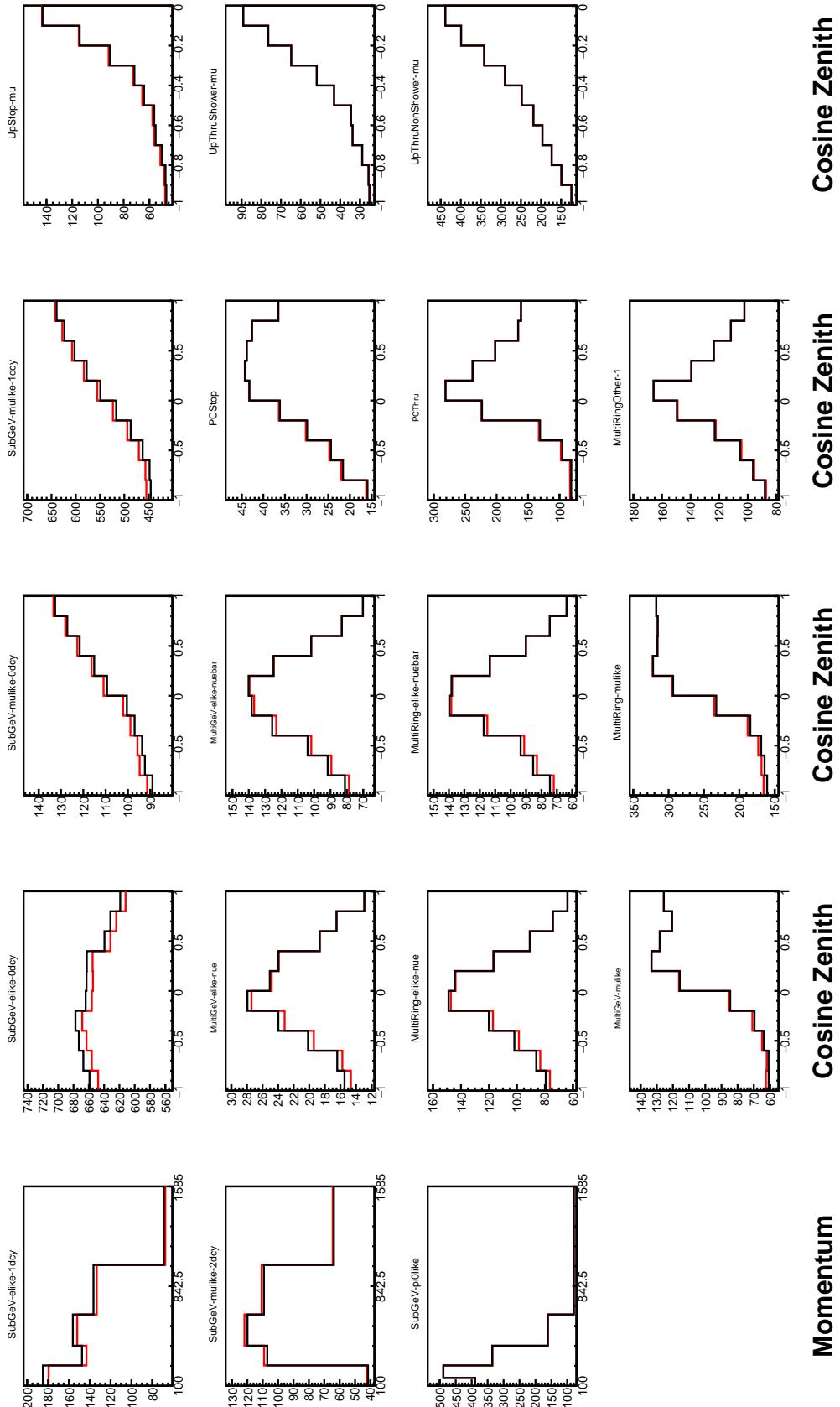
**Table 6.3:** The fully contained multiGeV samples used within this oscillation analysis. Both the sample name and description are given.

1804     Neutrino and antineutrino events are then separated by a second likelihood  
 1805     method ( $\nu/\bar{\nu}$  separation) detailed in [59]. This uses the number of decay electrons,

1806 the number of reconstructed rings, and the event’s transverse momentum. The  
1807 last two parameters are used because higher-energy samples tend to have more  
1808 pions produced above the Cherenkov threshold which results in more rings  
1809 compared to an antineutrino interaction. Furthermore, the angular distribution  
1810 also tends to be more forward peaked in antineutrino interactions as compared  
1811 to neutrino interactions [85]. These FC multiGeV sample definitions are de-  
1812 tailed in Table 6.3.

1813 The PC and up- $\mu$  samples are split by the amount of energy deposited within  
1814 the outer detector, into “stopping” and “through-going” samples. If an event  
1815 leaves the detector, the energy it takes with it has to be estimated which increases  
1816 the systematic uncertainty compared to events entirely contained within the  
1817 inner detector. This estimation is particularly poor at high energies, thus the  
1818 up- $\mu$  through-going events are not binned in reconstructed momentum. The  
1819 through-going up- $\mu$  are further separated by the presence of any electromagnetic  
1820 showering in the event, as the assumption of non-showering muon does not give  
1821 reliable reconstruction for these types of events [54]. In total, 13 FC, 2 PC, and  
1822 3 up- $\mu$  atmospheric samples are included within this analysis.

1823 The atmospheric samples are binned in direct observables: reconstructed  
1824 lepton momentum and direction, as given by Table A.1. The distribution of  
1825 the reconstructed lepton momentum (for samples that only have one bin in  
1826 reconstructed zenith angle) and reconstructed direction for each atmospheric  
1827 sample used within this analysis is illustrated in Figure 6.2. The by-mode  
1828 breakdown of each of the atmospheric samples is given in Appendix A.



**Figure 6.2:** Comparison of the SK-IV atmospheric samples between predictions made with the CP-violating Asimov A (Black) and CP-conserving Asimov B (Red) oscillation parameter sets (given in Table 2.2). The subGeV samples CCRES and  $\pi^0$ -like samples are given in their reconstructed lepton momentum. All other samples are presented in their reconstructed zenith angle projection.

Sample	$CC\nu_e$	$CC\bar{\nu}_e$	$CC(\nu_\mu + \bar{\nu}_\mu)$	$CC(\nu_\tau + \bar{\nu}_\tau)$	NC
SubGeV- <i>elike</i> -0dcy	72.17	23.3	0.724	0.033	3.77
SubGeV- <i>elike</i> -1dcy	86.81	1.773	7.002	0.062	4.351
SubGeV- <i>mulike</i> -0dcy	1.003	0.380	90.07	0.036	8.511
SubGeV- <i>mulike</i> -1dcy	0.023	0.	98.46	0.029	1.484
SubGeV- <i>mulike</i> -2dcy	0.012	0.	99.25	0.030	0.711
SubGeV- <i>pi0like</i>	6.923	2.368	0.928	0.011	89.77
MultiGeV- <i>elike</i> -nue	78.18	7.041	3.439	1.886	9.451
MultiGeV- <i>elike</i> -nuebar	56.68	37.81	0.174	0.614	4.718
MultiGeV- <i>mulike</i>	0.024	0.005	99.67	0.245	0.058
MultiRing- <i>elike</i> -nue	59.32	12.39	4.906	3.385	20
MultiRing- <i>elike</i> -nuebar	52.39	31.03	1.854	1.585	13.14
MultiRing- <i>mulike</i>	0.673	0.080	97.33	0.342	1.578
MultiRingOther-1	27.98	2.366	34.93	4.946	29.78
PCStop	8.216	3.118	84.45	0.	4.214
PCThru	0.564	0.207	98.65	0.	0.576
UpStop-mu	0.829	0.370	98.51	0.	0.289
UpThruNonShower-mu	0.206	0.073	99.62	0.	0.103
UpThruShower-mu	0.128	0.054	99.69	0.	0.132

**Table 6.4:** The purity of each atmospheric sample used within this analysis, broken down by charged current (CC) and neutral current (NC) interactions and which neutrino flavour interacted within the detector. Each row sums to 100% by definition. Asimov A oscillation parameter sets are assumed (given in Table 2.2). Electron neutrino and antineutrino events are separated to illustrate the ability of the separation likelihood cuts used within the multiGeV and multiring sample selections.

## 1829 6.2 Near Detector Beam Samples

1830 The near detector sample selections are documented in detail within [190] and  
1831 summarised below. Samples are selected based upon which of the two Fine  
1832 Grained Detector (FGD) the vertex is reconstructed in as well as the operating  
1833 mode of the beam: FHC or RHC. Wrong-sign neutrino background samples are  
1834 considered in the RHC mode in order to add additional constraints on model  
1835 parameters. Samples from the wrong-sign component of the FHC beam mode  
1836 are not included as they are statistically insignificant compared to those samples  
1837 already listed.

1838 The reconstruction algorithm uses a clustering algorithm to group hits within  
1839 the TPC. It then adds information from the upstream FGD to form a track  
1840 that passes through both sub-detectors. In FHC(RHC), the highest momentum  
1841 negative(positive) curvature track is defined as the muon candidate. Before  
1842 being assigned a sample, these candidate muon events must pass CC-inclusive  
1843 cuts, as defined in [191]:

- 1844 • Event Timing: The DAQ must be operational and the event must occur  
1845 within the expected beam time window consistent with the beam spill
- 1846 • TPC Requirement: The muon-candidate track path must intercept one or  
1847 more TPCs
- 1848 • Fiducial volume: The event must originate from within the fiducial volume  
1849 defined in [192]
- 1850 • Upstream Background: Remove events that have muon tracks that originate  
1851 upstream of the FGDs by requiring no high-momentum tracks within  
1852 150mm upstream of the candidate vertex. Additionally, events that occur  
1853 within the downstream FGD are vetoed if a secondary track starts within  
1854 the upstream FGD

- 1855     • Broken track removal: All candidates where the muon candidate is broken  
1856       in two are removed

- 1857     • Muon PID: Measurements of  $dE/dx$  in a TPC are used to distinguish muon-  
1858       like events, from electron-like or proton-like, using a likelihood cut

1859     In addition to these cuts, RHC neutrino events also have to undergo the  
1860       following cuts to aid in the separation of neutrino and antineutrino [193]:

- 1861     • TPC Requirement: The track path must intercept TPC2  
1862     • Positive Track: The highest momentum track must have a positive recon-  
1863       structed charge  
1864     • TPC1 Veto: Remove any events originating upstream of TPC1

1865     Once all CC-inclusive events have been determined, they are further split  
1866       by pion multiplicity: CC0 $\pi$ , CC1 $\pi$ , and CCOther. This breakdown targets the  
1867       specific interaction modes CCQE, CCRES, and other CC background interactions,  
1868       respectively. Pions in the TPCs and FGDs are selected by requiring a second track  
1869       to be observed, which is separate from the muon track and is in the same beam  
1870       spill window and sub-detector. If the pion originated within a FGD, it must also  
1871       pass through the sequential downstream TPC (TPC2 for FGD1, TPC3 for FGD2).

1872     CC0 $\pi$ , CC1 $\pi$ , and CCOther samples are defined with the following cuts:

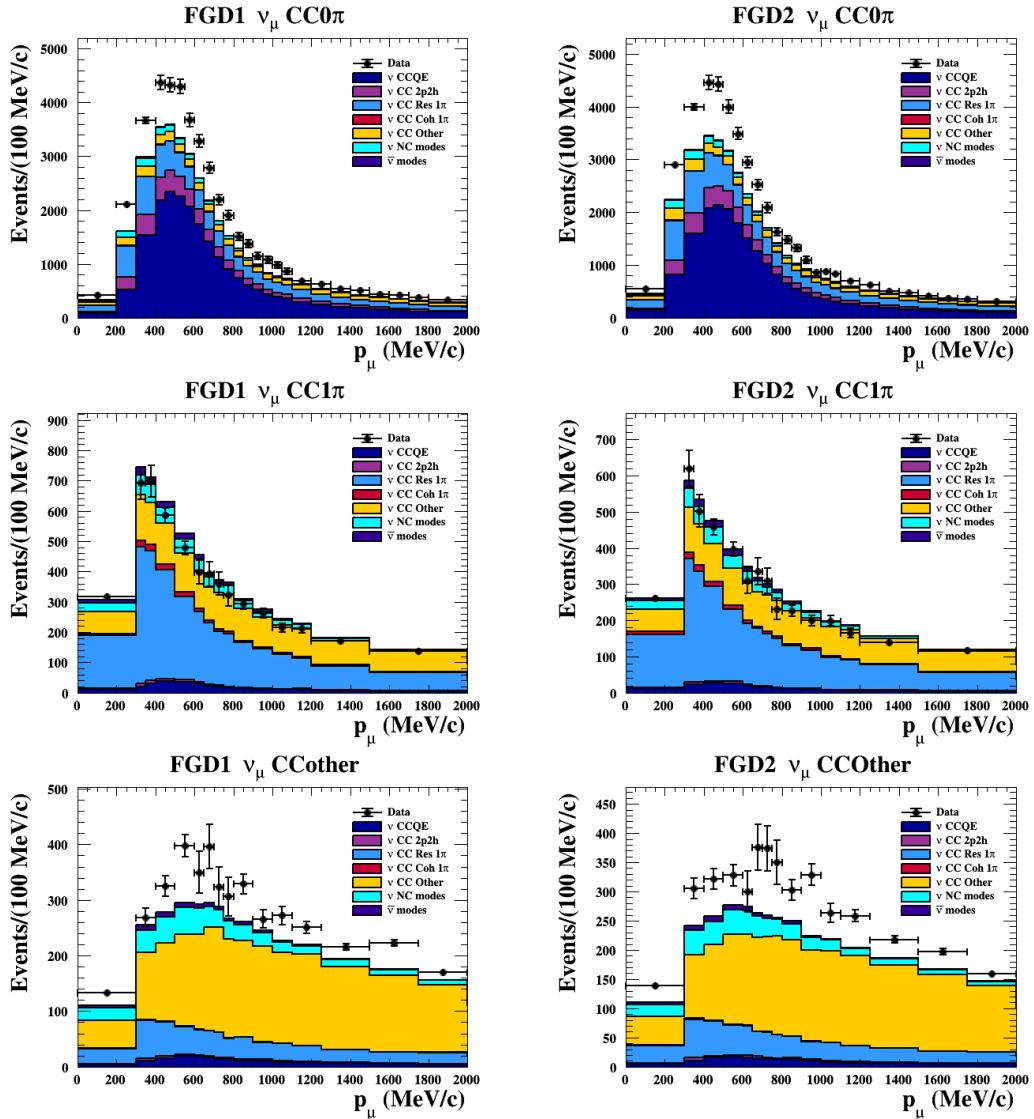
1873     **DB: Understand pion cuts at ND**

- 1874     •  $\nu_\mu$ CC0 $\pi$  **Selection:** No electrons in TPC and no charged pions or decay  
1875       electrons within the TPC or FGD

- 1876     •  $\nu_\mu$ CC1 $\pi$  **Selection:** Exactly one charged pion in either the TPC or FGD,  
1877       where the number of charged pions in the FGD is equal to the number of  
1878       decay electrons

- 1879     •  $\nu_\mu$ CCOther **Selection:** All events which are not classified into the above  
1880       two selections

Counting the three selections for each FGD in FHC and RHC running, including the wrong-sign background in RHC, 18 near detector samples are used within this analysis. These samples are binned in reconstructed lepton momentum (illustrated in Figure 6.3) and direction with respect to the beam. The binning is chosen such that each event has at least 20 Monte Carlo events in each bin [192]. This is to ensure that the bins are coarse enough to ensure the reduction of statistical errors, whilst also being fine enough to sample the high-resolution peak regions. The exact binning is detailed in [192].



**Figure 6.3:** The nominal Monte Carlo predictions compared to data for the FGD1 and FGD2 samples in neutrino beam mode, broken down into the  $CC\nu_\mu 0\pi$ ,  $CC\nu_\mu 1\pi$  and  $CC\nu_\mu$  Other categories. Figures taken from [190].

### **1889 6.3 Far Detector Beam Samples**

1890 The beam neutrino events which occur at the SK detector, which pass the  
 1891 reduction cuts detailed in section 5.3, are separated based on whether the beam  
 1892 was operating in FHC or RHC mode. The events are then separated into three  
 1893 samples: electron-like ( $1Re$ ), muon-like ( $1R\mu$ ), and  $CC1\pi^+$ -like ( $1Re1de$ ) which  
 1894 are observed as electron-like events with an associated decay electron [182].  
 1895 As discussed in section 6.1, positively charged pions emitted from neutrino  
 1896 interactions are more likely to produce decay electrons than negatively charged  
 1897 pions. Consequently, the  $CC1\pi^+$ -like sample is only selected when the beam is  
 1898 operating in FHC mode. Therefore, five beam samples measured at SK are  
 1899 used in this analysis.

1900 The fiducial volume definition for beam samples is slightly different from that  
 1901 used for the atmospheric samples. It uses both the distance to the closest wall  
 1902 (`dWall`) and the distance to the wall along the trajectory of the particle (`toWall`).  
 1903 This allows events that originate close to the wall but are facing into the tank to be  
 1904 included within the analysis, which would have otherwise been removed. These  
 1905 additional events are beneficial for a statistics-limited experiment. The exact  
 1906 cut values for both `dWall` and `toWall` are different for each of the three types of  
 1907 sample and are optimised based on T2K sensitivity to  $\delta_{CP}$  [180, 194]. They are:

1908 **1Re event selection** For an event to be classified as a  $1Re$ -like, the event must sat-  
 1909 isfy:

- 1910 • Fully-contained and have  $dWall > 80\text{cm}$  and  $toWall > 170\text{cm}$
- 1911 • Total of one ring which is reconstructed as electron-like with reconstructed  
 1912 momentum  $P_e > 100\text{MeV}$
- 1913 • Zero decay electrons are associated with the event
- 1914 • Passes  $\pi^0$  rejection cut discussed in section 5.2

1915      The zero decay electron cut removes non-CCQE interactions and the  $\pi^0$   
 1916     rejection cut is designed to remove neutral current  $\pi^0$  background events which  
 1917     can be easily reconstructed as 1Re-like events.

1918      The zero decay electron cut removes non-CCQE interactions and the  $\pi^0$   
 1919     rejection cut is designed to remove neutral current  $\pi^0$  background events which  
 1920     can be easily reconstructed as 1Re-like events.

1921   **CC1 $\pi^+$  event selection** This event selection is very similar to that of the 1Re  
 1922   sample. The only differences are that the `dWall` and `toWall` criteria are changed  
 1923   to  $> 50\text{cm}$  and  $> 270\text{cm}$ , respectively, and exactly one decay electron is required  
 1924   from the  $\pi^+$  decay.

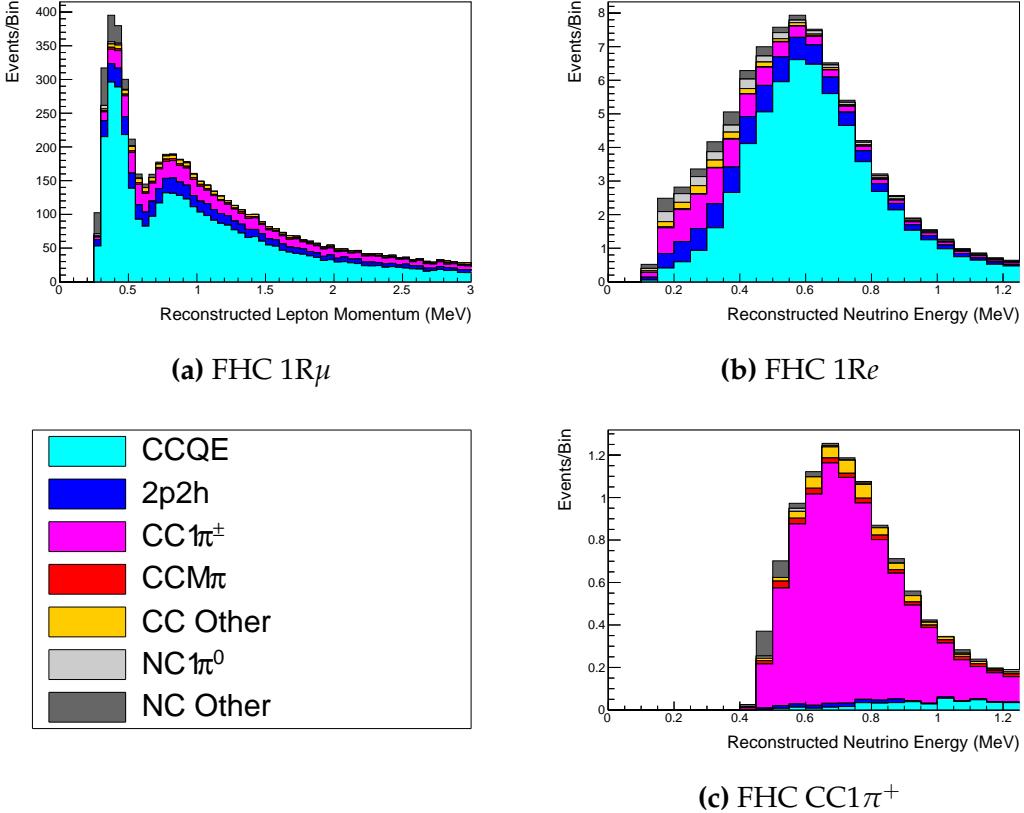
1925   **1R $\mu$  event selection** A 1R $\mu$ -like event is determined by the following cuts:

- 1926     • Fully-contained and have `dWall`  $> 50\text{cm}$  and `toWall`  $> 250\text{cm}$
- 1927     • Total of one ring which is reconstructed as muon-like with reconstructed  
   1928       momentum  $P_\mu > 200\text{MeV}$
- 1929     • Fewer than two decay electrons are associated with the event
- 1930     • Passes  $\pi^+$  rejection cut discussed in section 5.2

1931   All of these samples are binned in reconstructed neutrino energy. This is  
 1932   possible under a particular interaction mode assumption, as the direction from  
 1933   the source is known extremely well. For the 1Re-like and 1R $\mu$ -like samples,

$$E_\nu^{rec} = \frac{(M_N - V_{nuc})E_l - m_l^2/2 + M_N V_{nuc} - V_{nuc}^2/2 + (M_P^2 + M_N^2)/2}{M_N - V_{nuc} - E_l + P_l \cos(\theta_{beam})}. \quad (6.2)$$

1934   Where  $M_N$ ,  $M_P$  and  $m_l$  are the masses of the neutron, proton and outgoing  
 1935   lepton, respectively.  $V_{nuc} = 27\text{MeV}$  is the binding energy of the oxygen nucleus  
 1936   [182],  $\theta_{beam}$  is the angle between the beam and the direction of the outgoing  
 1937   lepton, and  $E_l$  and  $P_l$  are the energy and momentum of that outgoing lepton.



**Figure 6.4:** The reconstructed neutrino energy, as defined by Equation 6.2 and Equation 6.3, for the 1R $\mu$ -like, 1Re-like, and CC1 $\pi^+$ -like samples. The AsimovA oscillation parameters are assumed (given in Table 2.2). These samples are the FHC mode samples. For ease of viewing, the 1R $\mu$  sample only shows the  $0 \leq E_\nu^{rec} < 3.0\text{GeV}$  but the binning extends to 30.0GeV.

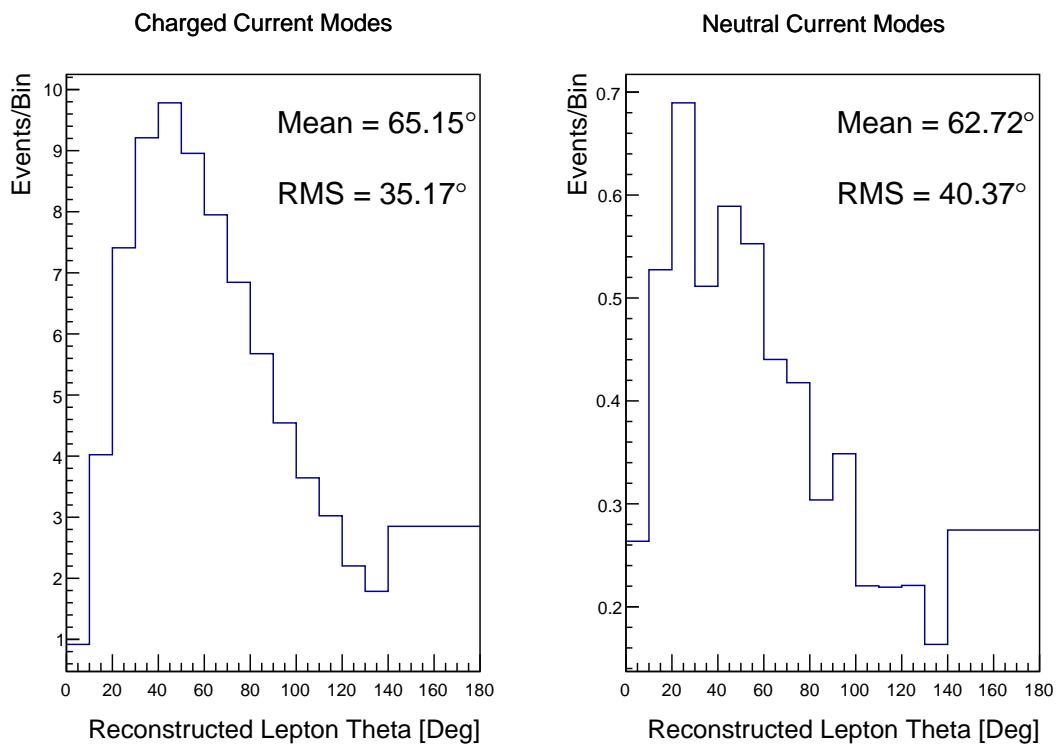
1938      The reconstructed neutrino energy of the CC1 $\pi^+$ -like events also accounts  
 1939      for the delta resonance produced within the interaction,

$$E_\nu^{rec} = \frac{2M_N E_l + M_{\Delta^{++}}^2 - M_N^2 - m_l^2}{2(M_N - E_l + P_l \cos(\theta_{beam}))}. \quad (6.3)$$

1940      Where  $M_{\Delta^{++}}$  is the mass of the delta baryon. Binding energy effects are not  
 1941      considered as a two-body process, with the delta baryon, is assumed. This follows  
 1942      the T2K oscillation analysis presented in [74], although recent developments of  
 1943      the interaction model in the latest T2K oscillation analysis do include effects  
 1944      from binding energy in this calculation [195].

1945      The reconstructed neutrino energy for the FHC samples is illustrated in  
 1946      Figure 6.4. As expected, the 1R $\mu$ -like and 1Re-like samples are heavily dominated

by CCQE interactions, with smaller contributions from 2p2h meson exchange and resonant pion production interactions. The CC1 $\pi^+$ -like sample predominantly consists of charged current resonant pion production interactions. The 1Re-like and CC1 $\pi^+$ -like samples are also binned by the angle between the neutrino beam and the reconstructed lepton momentum. This is to aid in charged current and neutral current separation, as indicated in Figure 6.5. This is because the neutral current backgrounds are predominantly due to  $\pi^0$ -decays, which decay into two  $\gamma$  rays. The opening angle of which (alongside the different final state kinematics) can produce a slightly broader angular distribution compared to the final state particles originating from charged current  $\nu_e$  interactions.



**Figure 6.5:** The distribution of the angle between the neutrino beam direction and the reconstructed final state lepton, for the FHC 1Re-like sample. The distribution is broken down by neutrino interaction mode into charged current (left) and neutral current (right) components. Asimov A oscillation parameter sets are assumed (given in Table 2.2). The RMS of the charged and neutral current plots are  $35.17^\circ$  and  $40.37^\circ$ , respectively.

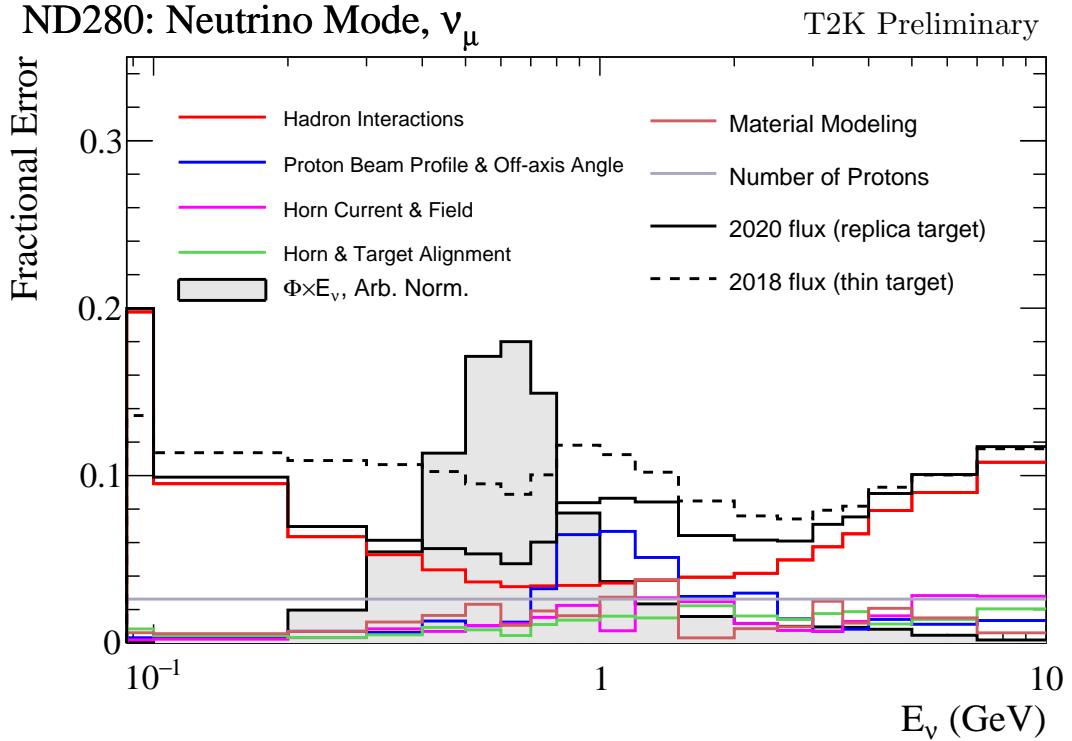
## 1957 6.4 Systematic Uncertainties

1958 The systematic model parameters for this analysis are split into groups, or blocks,  
1959 depending on their purpose. They consist of flux uncertainties, neutrino-matter  
1960 interaction systematics, and detector efficiencies. There are also uncertainties on  
1961 the oscillation parameters to which this analysis is not sensitive, namely  $\Delta m_{21}^2$   
1962 and  $\sin^2(\theta_{12})$ . These oscillation parameter uncertainties are taken from the 2020  
1963 PDG measurements [75]. As described in chapter 4, each model parameter used  
1964 within this analysis requires a prior uncertainty. This is provided via separate  
1965 covariance matrices for each block. The covariance matrices can include prior  
1966 correlations between parameters within a single block, but the separate treatment  
1967 means prior correlations can not be included for parameters in different groups.  
1968 Some parameters in these models have no reasonably motivated uncertainties  
1969 and are assigned flat priors which do not modify the likelihood penalty. In  
1970 practice, these flat prior parameters are actually assigned a Gaussian with a  
1971 very large width to ensure the covariance matrix is positive definite. They are  
1972 then checked at run time to determine if they contribute to the likelihood. The  
1973 flux, neutrino interaction, and detector modeling simulations have already been  
1974 discussed in section 5.1 and section 5.2. The uncertainties invoked within each  
1975 of these models are described below.

### 1976 6.4.1 Beam Flux

1977 The neutrino beam flux systematics are based upon the uncertainty in the mod-  
1978 eling of the components of the beam simulation. This includes the model of  
1979 hadron productions and reinteractions, the shape, intensity, and alignment of  
1980 the beam with respect to the target, and the uniformity of the magnetic field  
1981 produced by the horn, alongside other effects. The uncertainty, as a function  
1982 of neutrino energy, is illustrated in Figure 6.6 which includes a depiction of  
1983 the total uncertainty as well as the contribution from individual components.  
1984 The uncertainty around the peak of the energy distribution ( $E_\nu \sim 0.6\text{GeV}$ ) is

1985 dominated by uncertainties in the beam profile and alignment. Outside of this  
1986 region, uncertainties on hadron production dominate the error.



**Figure 6.6:** The total uncertainty evaluated on the near detector  $\nu_\mu$  flux prediction constrained by the replica-target data, illustrated as a function of neutrino energy. The solid(dashed) line indicates the uncertainty used within this analysis(the T2K 2018 analysis [196]). The solid histogram indicates the neutrino flux as a function of energy. Figure taken from [197].

1987 The beam flux uncertainties are described by one hundred parameters. They  
1988 are split between the ND280 and SK detectors and binned by neutrino flavour:  
1989  $\nu_\mu$ ,  $\bar{\nu}_\mu$ ,  $\nu_e$  and  $\bar{\nu}_e$ . The response is then broken down as a function of neutrino  
1990 energy. The bin density in the neutrino energy is the same for the  $\nu_\mu$  in FHC  
1991 and  $\bar{\nu}_\mu$  in RHC beams, and narrows for neutrino energies close to the oscillation  
1992 maximum of  $E_\nu = 0.6\text{GeV}$ . This binning is specified in Table 6.5. All of these  
1993 systematic uncertainties are applied as normalisation parameters with Gaussian  
1994 priors centered at 1.0 and error specified from a covariance matrix provided  
1995 by the T2K beam group [197].

Neutrino Flavour	Sign	Neutrino Energy Bin Edges (GeV)
$\mu$	Right	0., 0.4, 0.5, 0.6, 0.7, 1., 1.5, 2.5, 3.5, 5., 7., 30.
$\mu$	Wrong	0., 0.7, 1., 1.5, 2.5, 30.
$e$	Right	0., 0.5, 0.7, 0.8, 1.5, 2.5, 4., 30.
$e$	Wrong	0., 2.5, 30.

**Table 6.5:** The neutrino energy binning for the different neutrino flavours. “Right” sign indicates neutrinos in the FHC beam and antineutrinos in the RHC beam. “Wrong” sign indicates antineutrinos in the FHC beam and neutrinos in the RHC beam. The binning of the detector response is identical for the FHC and RHC modes as well as at ND280 and SK.

#### 1996 6.4.2 Atmospheric Flux

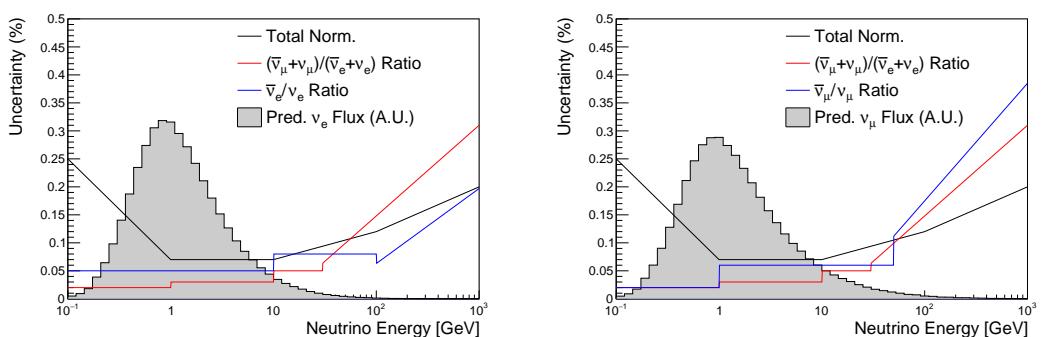
1997 The atmospheric neutrino flux is modeled by the HKKM model [50]. 16 systematic  
 1998 uncertainties are applied to control the normalisation of each neutrino flavour,  
 1999 energy, and direction. They are summarised below:

- 2000 • **Absolute Normalisation:** The overall normalisation of each neutrino flavour  
 2001 is controlled by two independent systematic uncertainties, for  $E_\nu < 1\text{GeV}$   
 2002 and  $E_\nu > 1\text{GeV}$ , respectively. This is driven mostly by hadronic interaction  
 2003 uncertainties for the production of pions and kaons [50]. The strength of  
 2004 the response is dependent upon the neutrino energy. The uncertainty is  
 2005 parameterized following Figure 11 in [50].
- 2006 • **Relative Normalisation:** Uncertainties on the ratio of  $(\nu_\mu + \bar{\nu}_\mu) / (\nu_e + \bar{\nu}_e)$   
 2007 are controlled by the difference between the HKKM model [50], FLUKA  
 2008 [53] and Bartol models [49]. Three independent parameters are applied in  
 2009 the energy ranges:  $E_\nu < 1\text{GeV}$ ,  $1\text{GeV} < E_\nu < 10\text{GeV}$ , and  $E_\nu > 10\text{GeV}$ .
- 2010 •  **$\nu/\bar{\nu}$  Normalisation:** The uncertainties in the  $\pi^+/\pi^-$  (and kaon equivalent)  
 2011 production uncertainties in the flux of  $\nu/\bar{\nu}$ . The response is applied using  
 2012 the same methodology as the relative normalisation parameters.
- 2013 • **Up/Down and Vertical/Horizontal Ratio:** Similar to the above two sys-  
 2014 tematics, the difference between the HKKM, FLUKA, and Bartol model

2015 predictions, as a function of  $\cos(\theta_Z)$ , is used to control the normalisation of  
 2016 events as a function of zenith angle.

- 2017 •  **$K/\pi$  Ratio:** Higher energy neutrinos ( $E_\nu > 10\text{GeV}$ ) mostly originate in  
 2018 kaon decay. Measurements of the ratio of  $K/\pi$  production [198] are used to  
 2019 control the systematic uncertainty of the expected ratio of pion and kaon  
 2020 production.
- 2021 • **Solar Activity:** As the 11-year solar cycle can affect the Earth's magnetic  
 2022 field, the flux of primary cosmic rays varies across the same period. The  
 2023 uncertainty is calculated by taking a  $\pm 1$  year variation, equating to a 10%  
 2024 uncertainty for the SK-IV period.
- 2025 • **Atmospheric Density:** The height of the interaction of the primary cosmic  
 2026 rays is dependent upon the atmospheric density. The HKKM assumes the  
 2027 US standard 1976 [150] profile. This systematic controls the uncertainty in  
 2028 that model.

2029 The total uncertainty is dominated by the absolute and relative normalisation  
 2030 parameters. The effect of which is illustrated in Figure 6.7. Generally, the  
 2031 uncertainty is large at low energy, reducing to  $O(10\%)$  around the peak of the  
 2032 flux distribution and then increasing once the neutrino energy exceeds 10GeV.



**Figure 6.7:** The uncertainty evaluated on the atmospheric  $\nu_e$  (left) and  $\nu_\mu$  (right) flux predictions. The absolute normalisation and flavour ratio uncertainties are given. The solid histogram indicates the neutrino flux as a function of energy.

2033      Updates to the HKKM and Bartol models are underway [155] to use a similar  
2034      tuning technique to that used in the beam flux predictions. After those updates,  
2035      it may be possible to include correlations in the hadron production uncertainty  
2036      systematics for beam and atmospheric flux predictions.

### 2037      6.4.3 Neutrino Interaction

2038      Neutrino interactions in the detectors are modeled by NEUT. The two indepen-  
2039      dent oscillation analyses, T2K-only [199] and the SK-only [59], have developed  
2040      separate interaction models. To maximise sensitivity out of this simultaneous  
2041      beam and atmospheric analysis, a correlated interaction model has been defined  
2042      in [200]. Where applicable, correlations allow the systematic uncertainties applied  
2043      to the atmospheric samples to be constrained by near detector neutrino beam  
2044      measurements. This can lead to stronger sensitivity to oscillation parameters  
2045      as compared to an uncorrelated model.

2046      The low-energy T2K systematic model has a more sophisticated treatment  
2047      of CCQE, 2p2h, and CCRES uncertainties, where extensive comparisons of  
2048      this model have been performed to external data [199]. However, the model  
2049      is not designed for high-energy atmospheric events, like those illustrated in  
2050      Figure 5.11. Therefore the high energy systematic model from the SK-only  
2051      analysis is implemented for the relevant multi-GeV, PC, and up- $\mu$  samples.  
2052      The T2K CCQE model is more sophisticated so it has been implemented for  
2053      all samples within this analysis, where separate low-energy and high-energy  
2054      dials have been implemented. The low-energy dials are constrained by the near  
2055      detector measurements and are uncorrelated to their high-energy counterparts.  
2056      The author of this thesis was responsible for implementing and validating the  
2057      combined cross-section model as documented in [200, 201].

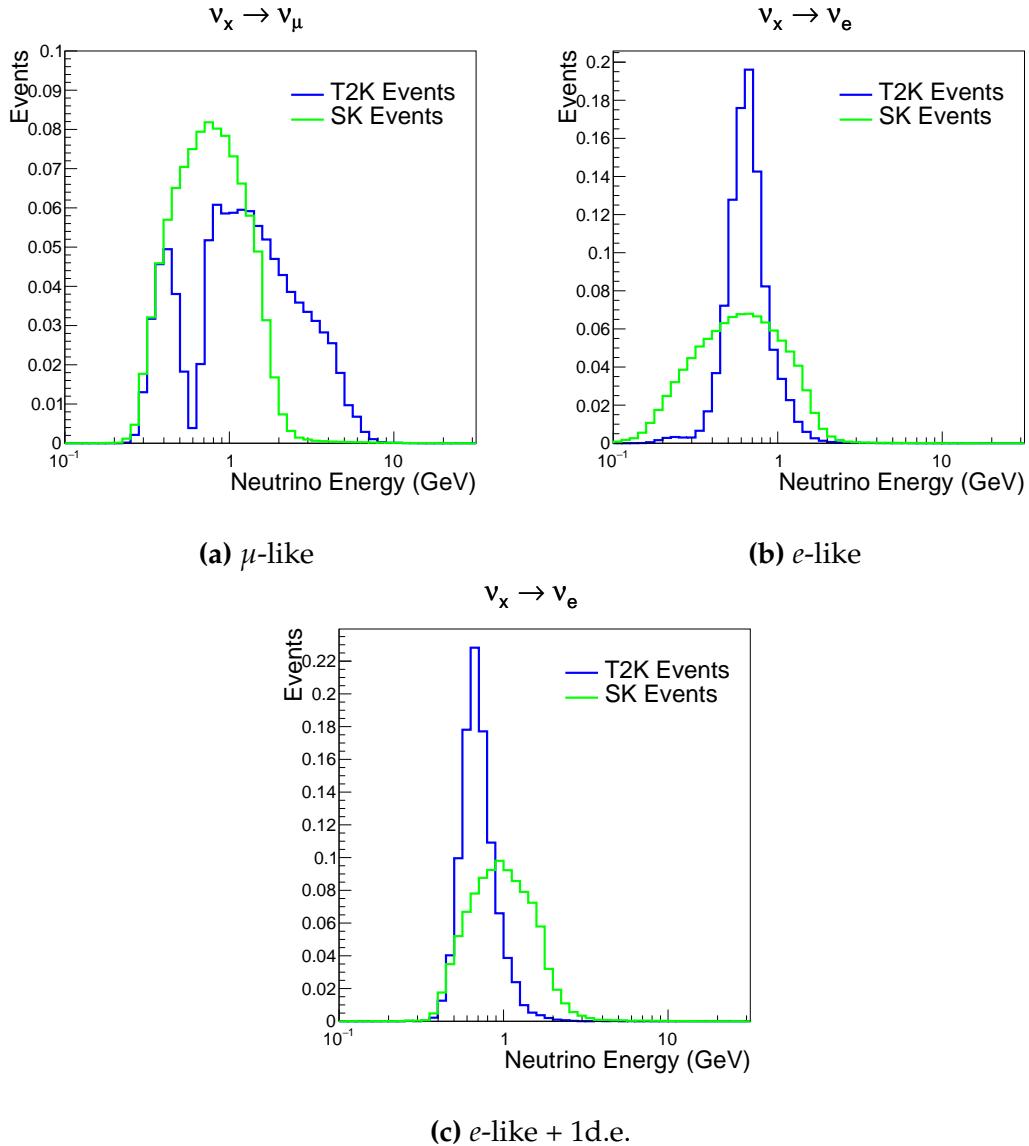
2058      The high energy systematic model includes parameters developed from  
2059      comparisons of Nieves and Rein-Seghal models which affect resonant pion  
2060      producing interactions, comparisons of the GRV98 and CKMT models which  
2061      control DIS interactions, and hadron multiplicity measurements which modulate

the normalisation of multi-pion producing events. The uncertainty on the  $\nu_\tau$  cross-section is particularly large and is controlled by a 25% normalisation uncertainty. These uncertainties are applied via normalisation or shape parameters. The former linearly scales the weight of all affected Monte-Carlo events, whereas the latter can increase or decrease a particular event's weight depending on its neutrino energy and mode of interaction. The response of the shape parameters is defined by third-order polynomial splines which return a weight for a particular neutrino energy. To reduce computational resources for the far detector fit, the response is binned by neutrino energy and sample binning: lepton momentum and cosine zenith binning for atmospheric splined responses and reconstructed neutrino energy and direction binning for beam samples. In total, 17 normalisation and 15 shape parameters are included in the high-energy model within this analysis.

Figure 6.8 indicates the predicted neutrino energy distribution for both beam and subGeV atmospheric samples. There is clearly significant overlap in neutrino energy between the subGeV atmospheric and beam samples, allowing similar kinematics in the final state particles. Figure 6.9 illustrates the fractional contribution of the different interaction modes per sample.

Comparing beam and atmospheric samples which target CCQE interactions (S.G. e-like 0de, S.G.  $\mu$ -like [0,1]de, [FHC,RHC] 1R  $\mu$ -like and [FHC,RHC] 1R e-like samples), there is a very similar contribution of CCQE, CC 2p2h, and CC1 $\pi^\pm$  interactions. The samples which target CC1 $\pi^\pm$  interactions, (S.G. e-like 0de, S.G.  $\mu$ -like 2de and FHC 1R+1d.e e-like) also consist of very similar mode interactions.

As a consequence of the similarity in energy and mode contributions, correlating the systematic model between the beam and subGeV atmospheric samples ensures that this analysis attains the largest sensitivity to oscillation parameters while still ensuring neutrino interaction systematics are correctly accounted for. Due to its more sophisticated CCQE and 2p2h model, the T2K systematic model was chosen as the basis of the correlated model.



**Figure 6.8:** The predicted neutrino energy distribution for subGeV atmospheric and beam samples. FHC and RHC beam samples are summed together Asimov A oscillation parameters are assumed (given in Table 2.2). Beam and atmospheric samples with similar cuts are compared against one another.

2092        The T2K systematic model [199] is applied in a similar methodology to the  
 2093        SK model parameters. It consists of 19 shape parameters and 24 normalisation  
 2094        parameters. Four additional parameters, which model the uncertainty in the  
 2095        binding energy, are applied in a way to shift the momentum of the lepton emitted  
 2096        from a nucleus. This controls the uncertainty specified on the 27MeV binding  
 2097        energy assumed within Equation 6.2. The majority of these parameters are

	CC QE	CC 2p2h	CC $1\pi^\pm$	CC $M\pi$	CC Other	NC $1\pi^0$	NC $1\pi^\pm$	NC $M\pi$	NC Coh.	NC Other
FHC 1R+1d.e. e-like	<b>0.04</b>	0.02	<b>0.83</b>	0.03	0.04	0.01	0.01	0.01	0.00	0.01
RHC 1R e-like	<b>0.62</b>	0.12	0.11	0.01	0.02	0.06	0.01	0.01	0.01	0.04
FHC 1R e-like	<b>0.68</b>	0.12	0.10	0.00	0.02	0.04	0.01	0.00	0.00	0.02
RHC 1R $\mu$ -like	<b>0.62</b>	0.13	0.17	0.02	0.03	0.00	0.02	0.00	0.00	0.00
FHC 1R $\mu$ -like	<b>0.62</b>	0.12	0.16	0.02	0.03	0.00	0.03	0.00	0.00	0.00
S.G. $\pi^0$ -like	<b>0.05</b>	0.01	0.02	0.00	0.01	<b>0.68</b>	0.06	0.07	0.06	0.04
S.G. $\mu$ -like 2de	<b>0.04</b>	0.01	<b>0.80</b>	0.10	0.04	0.00	0.00	0.00	0.00	0.00
S.G. $\mu$ -like 1de	<b>0.72</b>	0.11	0.12	0.01	0.02	0.00	0.01	0.00	0.00	0.00
S.G. $\mu$ -like 0de	<b>0.68</b>	0.11	0.10	0.01	0.02	0.01	0.05	0.01	0.00	0.02
S.G. e-like 1de	<b>0.05</b>	0.01	<b>0.75</b>	0.10	0.05	0.00	0.01	0.02	0.00	0.01
S.G. e-like 0de	<b>0.73</b>	0.11	0.10	0.01	0.02	0.02	0.00	0.00	0.00	0.00

**Figure 6.9:** The interaction mode contribution of each sample given as a fraction of the total event rate in that sample. Asimov A oscillation parameters are assumed (given in Table 2.2). The Charged Current (CC) modes are broken into quasi-elastic (QE), 2p2h, resonant charged pion production ( $1\pi^\pm$ ), multi-pion production ( $M\pi$ ), and other interaction categories. Neutral Current (NC) interaction modes are given in interaction mode categories:  $\pi^0$  production, resonant charged pion production, multi-pion production, and others.

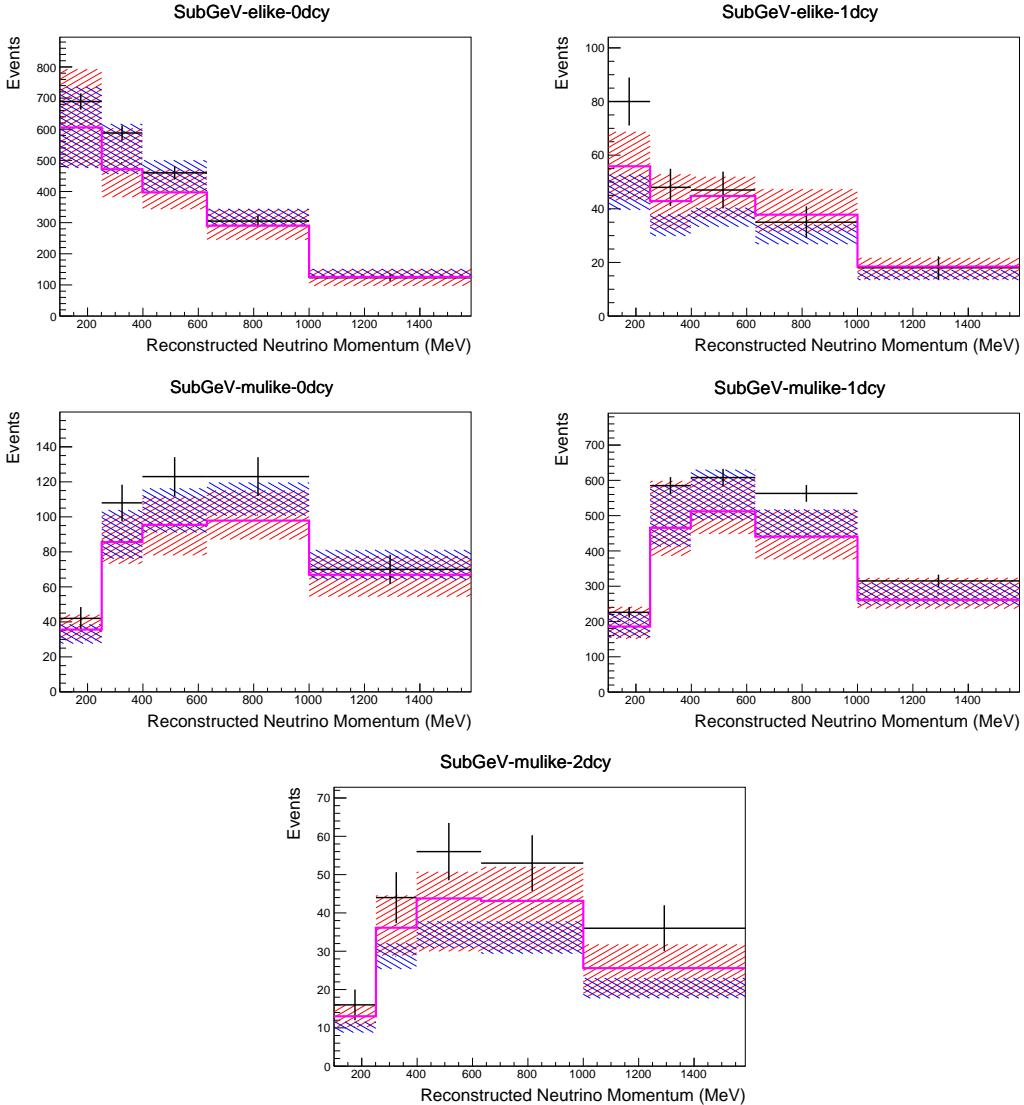
2098 assigned a Gaussian prior uncertainty. Those that have no reasonably motivated  
 2099 uncertainty, or those which have not been fit to external data, are assigned a  
 2100 flat prior which does not affect the penalty term.

2101 On top of the combination of the SK and T2K interaction models, several  
 2102 other parameters have been specifically developed for the joint oscillation anal-  
 2103 ysis. The majority of the atmospheric samples'  $\delta_{CP}$  sensitivity comes from the  
 2104 normalisation of subGeV electron-like events. These are modeled using a spectral  
 2105 function to approximate the nuclear ground state. However, the near detector is  
 2106 not able to constrain the model so an additional systematic is introduced which  
 2107 models an alternative Continuous Random Phase Approximation (CRPA) nuclear  
 2108 ground state. This dial approximates the event weights if a CRPA model had

been assumed rather than a spectral function. This dial only applies to  $\nu_e$  and  $\bar{\nu}_e$  as the near detector does not constraint  $\nu_e$  cross-section measurements. It is applied as a shape parameter.

Further additions to the model have been introduced due to the inclusion of the subGeV  $\pi^0$  atmospheric sample. This particularly targets charged current and neutral current  $\pi^0$  producing interactions to help constrain the systematic uncertainties. Therefore, an uncertainty that affects neutral current resonant  $\pi^0$  production is incorporated into this analysis. Comparisons of NEUT's NC resonant pion production predictions have been made to MiniBooNE [202] data and a consistent 16% to 21% underprediction is observed [200]. Consequently, a conservative 30% normalisation parameter is invoked.

Down-going events are mostly insensitive to oscillation parameters and can act similar to the near detector within an accelerator experiment (Details will be discussed in chapter 7). This region of phase space can act as a sideband and allows the cross-section model and near detector constraint to be studied. The distribution of events in this region is calculated using the technique outlined in subsection 4.3.4. The results are illustrated in Figure 6.10. For CCQE-targeting samples, the application of the near detector constraint is well within the statistical fluctuation of the down-going data. This means there is no significant tension is observed between the data and the Monte Carlo prediction after the near detector constraint is applied. This is not the case for samples with target CCRES interactions. The electron-like data is consistent with the constrained prediction at high reconstructed momenta but diverges at lower momentum, whereas the muon-like sample is under-predicted throughout the range of momenta. To combat this disagreement, an additional cross-section systematic dial, specifically designed to inflate the low pion momentum systematics was developed in [200]. This is a shape parameter implemented through a splined response.



**Figure 6.10:** Down-going atmospheric subGeV single-ring samples comparing the mean and error of the pre-fit and post-fit Monte Carlo predictions in red and blue, respectively. The magenta histogram illustrates the Monte Carlo prediction using the generated dial values. The black points illustrate the down-going data with statistical errors given. The mean and errors of the Monte Carlo predictions are calculated by the techniques documented in subsection 4.3.4. The pre-fit spectrum is calculated by throwing the cross-section and atmospheric flux dial values from the pre-fit covariance matrix. The post-fit spectrum is calculated by sampling the cross-section dial values from an ND fit MCMC chain, whilst still throwing the atmospheric flux dials from the pre-fit covariance.

#### 2136 6.4.4 Near Detector

2137 The systematics applied due to uncertainties arising from the response of the near  
 2138 detector is documented in [130]. The response is described by 574 normalisation  
 2139 parameters binned in the selected sample as well as momentum and angle,

2140  $P_\mu$  and  $\cos(\theta_\mu)$ , of the final-state muon. These are applied via a covariance  
2141 matrix with each parameter being assigned a Gaussian prior from that covariance  
2142 matrix. These normalisation parameters are built from underlying systematics,  
2143 e.g. pion secondary interaction systematics, which are randomly thrown and  
2144 the variation in each  $P_\mu \times \cos(\theta_\mu)$  bin is determined. Two thousand throws are  
2145 evaluated and a covariance matrix response is created. This allows significant  
2146 correlations between FGD1 and FGD2 samples, as well as adjacent  $P_\mu \times \cos(\theta_\mu)$   
2147 bins. Statistical uncertainties are accounted for by including fluctuations of each  
2148 event's weight from a Poisson distribution.

2149 Similar to the cross-section systematics, MaCh3 and BANFF are used to  
2150 constrain the uncertainty of these systematics through independent validations.  
2151 Each fitter generates a post-fit covariance matrix which is compared and passed  
2152 to the far-detector oscillation analysis working group. As the analysis presented  
2153 within this thesis uses the MaCh3 framework, a joint oscillation analysis fit of all  
2154 three sets of samples and their respective systematics is performed.

#### 2155 6.4.5 Far Detector

2156 Two configurations of the far detector systematic model implementation have  
2157 been considered. Firstly, the far detector systematic uncertainties for beam and  
2158 atmospheric samples are taken from their respective analysis inputs, denoted  
2159 “official inputs” analysis, with no correlations assumed between the beam and at-  
2160 mospheric samples. The beam- and atmospheric-specific inputs are documented  
2161 in subsubsection 6.4.5.1 and subsubsection 6.4.5.2. Secondly, an alternative  
2162 detector model has been developed which correlates the response of the SK  
2163 detector systematics between the beam and atmospheric samples. Here, the  
2164 distribution of parameters used for applying event cuts (e.g. electron-muon  
2165 PID separation) is modified within the fit. It follows a similar methodology to  
2166 the beam far detector systematics implementation but performs a joint fit of  
2167 the beam and atmospheric data. This alternative implementation is detailed  
2168 in subsubsection 6.4.5.3.

2169 **6.4.5.1 Beam Samples**

2170 There are 45 systematics which describe the response of the far detector to  
2171 beam events [182], split into 44 normalisation parameters and one energy scale  
2172 systematic. The energy scale systematic is applied as a multiplicative scaling  
2173 of the reconstructed neutrino energy. It is estimated from data-to-Monte Carlo  
2174 differences in the stopping muon sample in [184] and found to be 2.1%. The  
2175 normalisation parameters are assigned a Gaussian error centered at one with  
2176 width taken from a covariance matrix. A detailed breakdown of the generation  
2177 of the covariance matrix is found in [194]. To build the covariance matrix, a fit  
2178 is performed on atmospheric data which has been selected using beam sample  
2179 selection cuts. These cuts use the variables,  $L^i$ , where the index  $i$  is detailed in  
2180 Table 6.6. Each  $L^i$  is a smear,  $\alpha$ , and shift,  $\beta$  parameter such that,

$$L_j^i \rightarrow \bar{L}_j^i = \alpha_j^i L + \beta_j^i \quad (6.4)$$

2181 Where  $L_j^i$  ( $\bar{L}_j^i$ ) correspond to nominal(varied) PID cut parameters given in  
2182 Table 6.6. The shift and smear parameters are nuisance parameters with no prior  
2183 constraints. They are binned by final-state topology,  $j$ , where the binning is given  
2184 in Table 6.7. The final-state topology binning is because the detector will respond  
2185 differently to events that have one or multiple rings. For example, the detector  
2186 will be able to distinguish single-ring events better than two overlapping ring  
2187 events, resulting in different systematic uncertainty for one-ring events compared  
2188 to two-ring events. This approach is used to allow the cut parameter distributions  
2189 to be modified within the fit, allowing for better data to Monte Carlo agreement.

Cut Variable	Parameter Name
0	<code>fitQun e/mu PID</code>
1	<code>fitQun e/pi0 PID</code>
2	<code>fitQun mu/pi PID</code>
3	<code>fitQun Ring-Counting Parameter</code>

**Table 6.6:** List of cut variables that are included within the shift/smear fit documented in [194].

Category	Description
1e	Only one electron above Cherenkov threshold in the final state
1 $\mu$	Only one muon above Cherenkov threshold in the final state
1e+other	One electron and one or more other charged particles above Cherenkov threshold in the final state
1 $\mu$ +other	One muon and one or more other charged particles above Cherenkov threshold in the final state
1 $\pi^0$	Only one $\pi^0$ in the final state
1 $\pi^\pm$ or 1p	Only one hadron (typically charged pion or proton) in the final state
Other	Any other final state

**Table 6.7:** Reconstructed event topology categories on which the SK detector systematics [194] are based.

2190        The mis-modeling of  $\pi^0$  events is also considered. If one of the two rings  
 2191        from a  $\pi^0$  event is missed, this will be reconstructed as a CC $\nu_e$ -like event. This  
 2192        is one of the largest systematics hindering the electron neutrino appearance  
 2193        analyses. Consequently, additional systematics have been introduced to con-  
 2194        strain the mis-modeling of  $\pi^0$  events in SK, binned by reconstructed neutrino  
 2195        energy. To evaluate this systematic uncertainty, a set of “hybrid- $\pi^0$ ” samples is  
 2196        constructed. These events are built by overlaying one electron-like ring from  
 2197        the SK atmospheric neutrino samples or decay electron ring from a stopping  
 2198        cosmic ray muon with one simulated photon ring. Both rings are chosen so  
 2199        that momenta and opening angle follow the decay kinematics of NC  $\pi^0$  events  
 2200        from the T2K-MC. Hybrid- $\pi^0$  Monte Carlo samples with both rings from the  
 2201        SK Monte Carlo are produced to compare with the hybrid- $\pi^0$  data samples and  
 2202        the difference in the fraction of events that pass the  $\nu_e$  selection criteria is used  
 2203        to assign the systematic errors. In order to investigate any data to Monte Carlo  
 2204        differences that may originate from either the higher energy ring or lower energy  
 2205        ring, two samples are built; a sample in which the electron constitutes the higher  
 2206        energy ring from the  $\pi^0$  decay (called the primary sample) and another one in  
 2207        which it constitutes the lower energy ring (called the secondary sample). The  
 2208        standard T2K  $\nu_e$  fitQun event selection criteria are used to select events.

2209        Final contributions to the covariance matrix are determined by supplemen-  
 2210        tary uncertainties obtained by comparing stopping muon data to Monte Carlo

prediction, as first introduced in section 5.2. The efficiency of tagging decay electrons is estimated by the stopping muon data to Monte Carlo differences by comparing the number of one decay electron events to the number of events with one or fewer decay electrons. Similarly, the rate at which fake decay electrons are reconstructed by `fitQun` is estimated by comparing the number of two decay electron events to the number of events with one or two reconstructed decay electrons. The two sources of systematics are added in quadrature weighted by the number of events with one true decay electron yielding a 0.2% systematic uncertainty. A fiducial volume systematic of  $\pm 2.5\text{cm}$  which corresponds to a 0.5% shift in the normalisation of events is also applied. Additional normalisation uncertainties based on neutrino flavour and interaction mode are also defined in [182, 203, 204].

Two additional sources of uncertainty are included: secondary and photoneuclear interactions. These are estimated by varying the underlying parameters are building a distribution of sample event rates. These contributions are then added in quadrature to the above covariance matrix.

#### DB: Matrix Diagonals from SK Det Matrix

##### 6.4.5.2 Atmospheric Samples

The detector systematics for atmospheric samples, documented in [85], are split into two sub-groups: those which are related to particle identification and ring counting systematics, and those which are related to calibration, separation, and reduction uncertainties.

The particle identification systematics consist of five parameters. The ring separation systematic enforces an anti-correlated response between the single-ring and multi-ring samples. This is implemented as a fractional increase/decrease in the overall normalisation of each sample, depending on the distance to the nearest wall from an event's vertex. The coefficients of the normalisation are estimated prior to the fit and depend on the particular atmospheric sample. Two electron-muon separation systematics are included within this model which

2240 anti-correlates the response of the electron-like and muon-like samples: one for  
2241 single-ring events and another for multi-ring events.

2242 The multi-ring electron-like separation likelihood, discussed in section 6.1,  
2243 encodes the ability of the detector to separate neutrino from anti-neutrino events.  
2244 Two normalisation parameters vary the relative normalisation of multi-ring  $\nu_e$   
2245 and  $\bar{\nu}_e$  samples whilst keeping a consistent overall event rate.

2246 There are 22 systematics related to calibration measurements, including effects  
2247 from backgrounds, reduction, and showering effects. They are documented in  
2248 [85] and are briefly summarised in Table 6.8. They are applied via normalisation  
2249 parameters, with the separation systematics requiring the conservation of event  
2250 rate across all samples.

#### 2251 6.4.5.3 Correlated Detector Model

2252 A complete uncertainty model of the SK detector would be able to determine  
2253 the systematic shift on the sample spectra for a variation of the underlying  
2254 parameters, e.g. PMT angular acceptance. However, this is computationally  
2255 intensive, requiring Monte Carlo predictions to be made for each plausible  
2256 variation. Consequently, an effective parameter model has been utilised for  
2257 a correlated detector model following from the T2K-only model implementation  
2258 documented in subsubsection 6.4.5.1. It correlates the detector systematics  
2259 between the far-detector beam and subGeV atmospheric samples due to their  
2260 similar energies and interaction types. As there are no equivalent beam samples,  
2261 the multi-GeV, multiring, PC, and Up- $\mu$  samples will be subject to the particle  
2262 identification systematics implementation as described in subsubsection 6.4.5.2  
2263 rather than using this correlated detector model. The calibration systematics also  
2264 described in the aforementioned chapter still apply to all atmospheric samples.

2265 The correlated detector model utilises the same smear and shift parameters  
2266 documented in subsubsection 6.4.5.1, split by final state topology. Beyond this,  
2267 the shift and smear parameters are split by visible energy deposited within the  
2268 detector, with binning specified in Table 6.9. This is because atmospheric events

Index	Description
0	Partially contained reduction
1	Fully contained reduction
2	Separation of fully contained and partially contained events
3	Separation of stopping and through-going partially contained events in top of detector
4	Separation of stopping and through-going partially contained events in barrel of detector
5	Separation of stopping and through-going partially contained events in bottom of detector
6	Background due to cosmic rays
7	Background due to flasher events
8	Vertex systematic moving events into and out of fiducial volume
9	Upward going muon event reduction
10	Separation of stopping and through-going in upward going muon events
11	Energy systematic in upward going muon events
12	Reconstruction of the path length of upward going muon events
13	Separation of showering and non-showering upward going muon events
14	Background of stopping upward going muon events
15	Background of non-showering through-going upward going muon events
16	Background of showering through-going upward going muon events
17	Efficiency of tagging two rings from $\pi^0$ decay
18	Efficiency of decay electron tagging
19	Background from downgoing cosmic muons
20	Asymmetry of energy deposition in tank
21	Energy scale deposition

**Table 6.8:** Sources of systematic errors specified within the grouped into the “calibration” systematics model.

are categorised by subGeV and multi-GeV events based on visible energy, so this splitting is required when correlating the systematic model for beam and atmospheric events. Alongside the technical requirement, higher energy events will be better reconstructed due to fractionally less noise within the detector. As a result of the inclusion of visible energy binning, Equation 6.4 becomes

$$L_{jk}^i \rightarrow \bar{L}_{jk}^i = \alpha_{jk}^i L + \beta_{jk}^i, \quad (6.5)$$

where  $k$  is the visible energy bin.

The implementation of this systematic model takes the events reconstructed values of the cut parameters, modifies them by the particular shift and smear

Index	Range (MeV)
0	$30 \geq E_{vis} > 300$
1	$300 \geq E_{vis} > 700$
2	$700 \geq E_{vis} > 1330$
3	$E_{vis} \geq 1330$

**Table 6.9:** Visible energy binning for which the correlated SK detector systematics are based

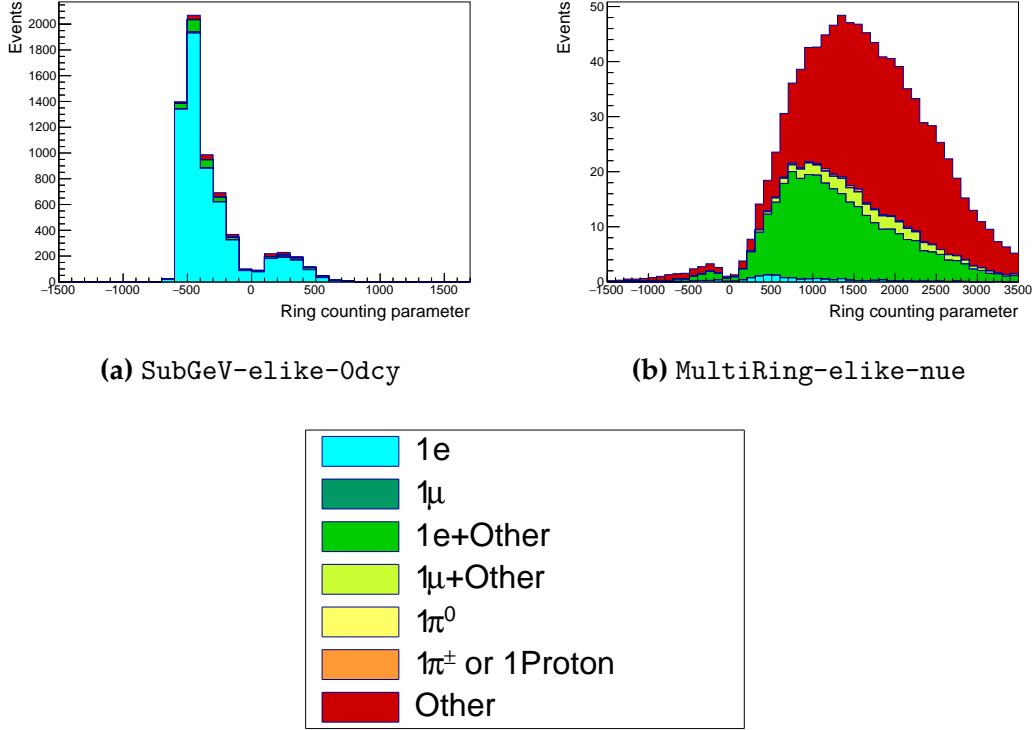
parameter for that event, and then re-applies event selection. This causes event migration, which is a new feature incorporated into the MaCh3 framework which is only achievable due to the event-by-event reweighting scheme.

Particular care has to be taken when varying the ring counting parameter. This is because the number of rings is a finite value (one-ring, two-ring, etc.) which can not be continuously varied through this shift and smear technique. Consequently a continuous ring counting parameter,  $RC_i$ , is calculated for the  $i^{th}$  event, following the definition in [181]: the preferred likelihoods from all considered one-ring ( $L_{1R}$ ) and two-ring ( $L_{2R}$ ) fits are determined. The difference is computed as  $\Delta_{LLH} = \log(L_{1R}) - \log(L_{2R})$ . The ring counting parameter is then defined as

$$RC_i = \text{sgn}(\Delta_{LLH}) \times \sqrt{|\Delta_{LLH}|}, \quad (6.6)$$

where  $\text{sgn}(x) = x/|x|$ . This ring counting parameter corresponds to an intermediate likelihood value used within the `fitQun` algorithm to decide the number of rings associated with a particular event. However, fake-ring merging algorithms are applied after this likelihood value is used. Consequently, this ring counting parameter does not always exactly correspond to the number of reconstructed rings. This can be seen in Figure 6.11.

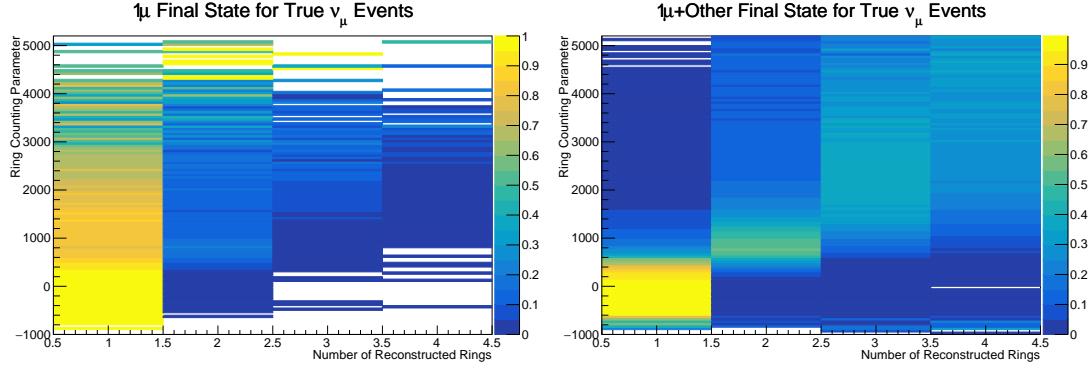
As the `fitQun` algorithm does not provide a likelihood value after the fake-ring algorithms have been applied, the ring counting parameter distribution is correlated to the final number of reconstructed rings through “maps”. These are two-dimensional distributions of the ring counting parameter and the final number of reconstructed rings. An example is illustrated in Figure 6.12. In



**Figure 6.11:** The ring counting parameter as defined in Equation 6.6 for the SubGeV-else-like-0dcy and MultiRing-else-like-nue samples.

principle, the `fitQun` reconstruction algorithm should be re-run after the variation in the ring counting parameter. However, this is not computationally viable. Therefore the “maps” are used as a reweighting template.

The maps are split by final state topology and true neutrino flavour and all `fitQun`-reconstructed Monte Carlo events are used to fill them. The maps are row-normalised to represent the probability of  $X$  rings for a given  $RC_i$  value. Prior to the oscillation fit, an event’s nominal weight is calculated as  $W^i(N_{Rings}^i, L_{jk}^i)$ , where  $N_{Rings}^i$  is the reconstructed number of rings for the  $i^{th}$  event and  $W^i(x, y)$  is the bin content in map associated with the  $i^{th}$  event, where  $x$  is the number of rings and  $y$  is the ring counting parameter. Then during the fit, the value of  $R = W^i(N_{Rings}^i, \bar{L}_{jk}^i) / W^i(N_{Rings}^i, L_{jk}^i)$  is calculated as the event weight for the  $i^{th}$  event. This is the only cut variable that uses a reweighting technique rather than event migration.

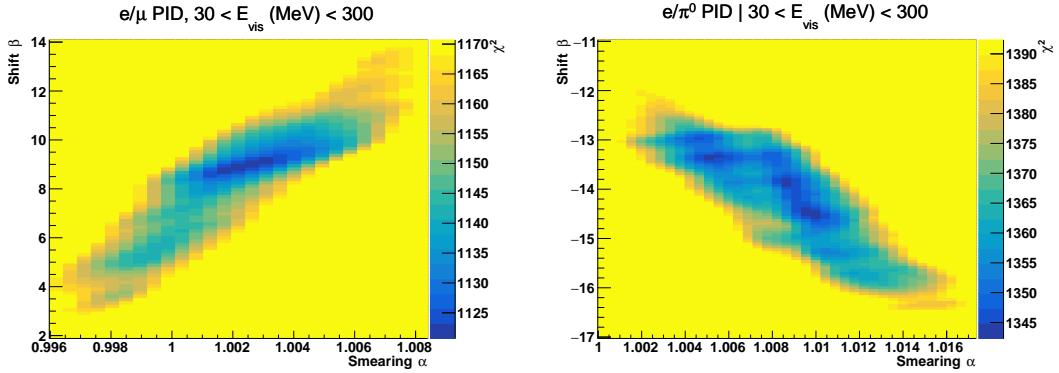


**Figure 6.12:** The ring counting parameter, defined in Equation 6.6, as a function of the number of reconstructed rings as found by the `fitTQu` reconstruction algorithm. Left: true  $\nu_\mu$  events with only one muon above the Cherenkov threshold in the final state. Right: true  $\nu_\mu$  events with one muon and at least one other charged particle above the Cherenkov threshold in the final state.

2312     The  $\pi^0$  systematics introduced in subsection 6.4.4 are applied via a covariance  
 2313     matrix. This is not possible in the alternative model as no covariance matrix  
 2314     is used. Thus, the implementation of the  $\pi^0$  systematics has been modified.  
 2315     The inputs from the hybrid  $\pi^0$  sample are included via the use of “ $\chi^2$  maps”,  
 2316     which are two-dimensional histograms in  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters over some  
 2317     range. Illustrative examples of the  $\chi^2$  maps are given in Figure 6.13. Due to  
 2318     their nature, the shift and smear parameters are typically very correlated. A  
 2319     map is produced for each cut parameter given in Table 6.6 and for each visible  
 2320     energy bin given in Table 6.9.

2321     The maps are filled through the  $\chi^2$  comparison of the hybrid  $\pi^0$  Monte Carlo  
 2322     and data in the particle identification parameters documented in Table 6.6. The  
 2323     Monte Carlo distribution is modified by the  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  scaling, whilst cross-  
 2324     section and flux nuisance parameters are thrown from their prior uncertainties.  
 2325     The  $\chi^2$  between the scaled Monte Carlo and data is calculated and the relevant  
 2326     point in the  $\chi^2$  map is filled.

2327     The implementation within this alternative detector model is to add the bin  
 2328     contents of the maps, for the relevant values of the  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters, to  
 2329     the likelihood penalty. Only 1 $\pi^0$  final state topology shift and smear parameters  
 2330     use this prior uncertainty.



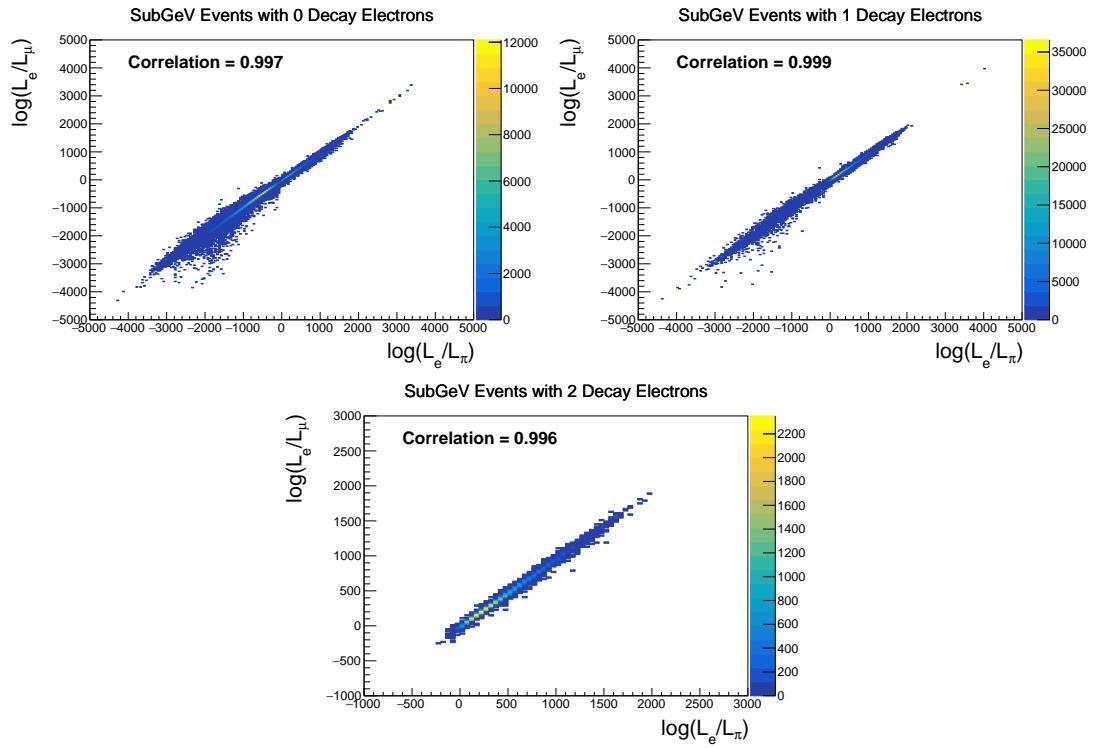
**Figure 6.13:** The  $\chi^2$  between the hybrid- $\pi^0$  Monte Carlo and data samples, as a function of smear ( $\alpha$ ) and shift ( $\beta$ ) parameters, for events which have  $1\pi^0$  final state topology. Left: Electron-muon separation PID parameter for events with  $30 \leq E_{\text{vis}}(\text{MeV}) < 300$ . Right: Electron- $\pi^0$  separation PID parameter for events with  $30 \leq E_{\text{vis}}(\text{MeV}) < 300$ .

Similarly, the implementation of the supplementary systematics documented in subsubsection 6.4.5.1 needs to be modified. A new framework [205] was built in tandem with the T2K-SK working group [182] so the additional parameters can be incorporated into the MaCh3 framework. These are applied as normalisation parameters, depending on the particular interaction mode, number of tagged decay electrons, and whether the primary particle generated Cherenkov light. They are assigned Gaussian uncertainties with widths described by a covariance matrix. Furthermore, the secondary interaction and photo-nuclear effects need to be accounted for in this detector model using a different implementation than that in subsubsection 6.4.5.1. This was done by including a shape parameter for each of the secondary interactions and the photo-nuclear systematic parameters.

There are a total of 224  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters, of which 32 have prior constraints from the hybrid  $\pi^0$  samples.

One final complexity of this correlated detector model is that the two sets of samples, beam and subGeV atmospheric, use slightly different parameters to distinguish electron and muon-like events. The T2K samples use the value of  $\log(L_e/L_\mu)$  whereas the atmospheric samples use the value of  $\log(L_e/L_\pi)$ , where  $L_X$  is the likelihood for hypothesis X. This is because the T2K fits use single-ring fiTQun fitting techniques, whereas multi-ring fits are applied to the atmospheric samples where only the electron and pion hypothesis are considered.

2351 The correlation between the two likelihood ratios is illustrated in Figure 6.14. As  
 2352 discussed in section 5.2, the pion hypothesis is a very good approximation of the  
 2353 muon hypothesis due to their similar mass. Consequently, using the same shift  
 2354 and smear parameters correlated between the beam and subGeV atmospheric  
 2355 samples is deemed a good approximation.



**Figure 6.14:** The distribution of  $\log(L_e/L_\mu)$  compared to  $\log(L_e/L_\pi)$  for subGeV events with zero (top left), one (top right) or two (bottom) decay electrons. The correlation in the distribution is calculated as 0.997, 0.999 and 0.996, respectively.

# 7

2356

2357

## Oscillation Probability Calculation

2358 It is important to understand how and where the sensitivity to the oscillation parameters comes from for both atmospheric and beam samples. An  
2359 overview of how these samples respond to changes in  $\delta_{CP}$ ,  $\Delta m_{32}^2$ , and  $\sin^2(\theta_{23})$   
2360 is given in section 7.1. This section also explains the additional complexities  
2361 involved when performing an atmospheric neutrino analysis as compared to  
2362 a beam-only analysis.

2364 Without additional techniques, atmospheric sub-GeV upward-going neutrinos ( $E_\nu < 1.33\text{GeV}, \cos(\theta_Z) < 0.$ ) can artificially inflate the sensitivity to  $\delta_{CP}$   
2365 due to the quickly varying oscillation probability in this region. Therefore, a  
2366 “sub-sampling” approach has been developed to reduce these biases ensuring  
2367 accurate and reliable sensitivity measurements. This technique ensures that small-  
2368 scale unresolvable features of the oscillation probability have been averaged over  
2369 whilst the large-scale features in the oscillation probability are unaffected. The  
2370 documentation and validation of this technique are found in section 7.2. The  
2371 oscillation probability calculation is computationally intensive due to the large  
2372 number of matrix multiplications needed. Consequently, the CUDAProb3 imple-  
2373 mentation choice made within the fitting framework, as detailed in section 7.3,  
2374 ensures that the analysis can be done in a timely manner.

Whilst the beam neutrinos are assumed to propagate through a constant density slab of material, the density variations through the Earth result in more complex oscillation patterns for atmospheric neutrinos. Furthermore, the uncertainty in the electron density can modify the oscillation probability for the denser core layers of the Earth. The model of the Earth used within this analysis is detailed in section 7.4. This includes information about the official SK-only methodology as well as improvements that have been made to remove some of the approximations used in that analysis. Another complexity of atmospheric neutrino oscillation studies is that the height of production in the atmosphere is not known on an event-by-event basis. An analytical averaging technique that approximates the uncertainty of the oscillation probability has been followed, with the author of this thesis being responsible for the implementation and validation. This implementation of an external technique is described in section 7.5.

## 7.1 Overview

**DB: Should this be moved into an earlier chapter? The selections chapter references the matter resonance which has not yet been explained at that point**

The analysis presented within this thesis focuses on the determination of oscillation parameters from atmospheric and beam neutrinos. Whilst subject to the same oscillation formalism, the way in which the two samples have sensitivity to the different oscillation parameters differs significantly.

Atmospheric neutrinos have a varying baseline, or “path length”  $L$ , such that the distance each neutrino travels before interacting is dependent upon the zenith angle,  $\theta_Z$ . As primary cosmic rays can interact anywhere between the Earth’s surface and  $\sim 50\text{km}$  above that, the height,  $h$ , in the atmosphere at which the neutrino was generated also affects the path length,

$$L = \sqrt{(R_E + h)^2 - R_E^2 (1 - \cos^2(\theta_Z))} - R_E \cos(\theta_Z). \quad (7.1)$$

2401     DB: Ask Giles what he means about the horizontal issue and Layering in  
 2402     the ProdH section

2403     Where  $R_E = 6,371\text{km}$  is the Earth's radius. Consequently, the oscillation  
 2404     probability is dependent upon two parameters,  $\cos(\theta_Z)$  and  $E_\nu$ .

2405     The oscillation probability used within this analysis is based on [21]. The  
 2406     neutrino wavefunction in the vacuum Hamiltonian evolves in each layer of  
 2407     constant matter density via

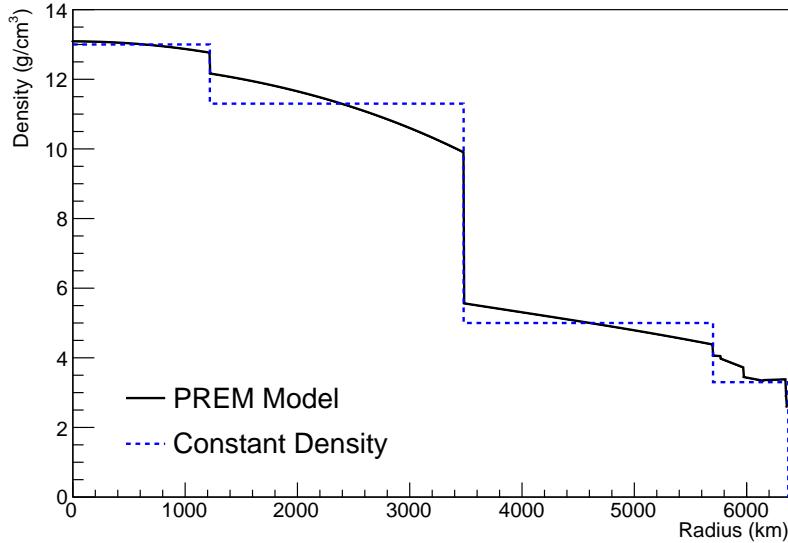
$$i \frac{d\psi_j(t)}{dt} = \frac{m_j^2}{2E_\nu} \psi_j(t) - \sum_k \sqrt{2} G_F N_e U_{ej} U_{ke}^\dagger \psi_k(t), \quad (7.2)$$

2408     where  $m_j^2$  is the square of the  $j^{th}$  vacuum eigenstate mass,  $E_\nu$  is the neutrino  
 2409     energy,  $G_F$  is Fermi's constant,  $N_e$  is the electron number density and  $U$  is the  
 2410     PMNS matrix. The transformation  $N_e \rightarrow -N_e$  and  $\delta_{CP} \rightarrow -\delta_{CP}$  is applied  
 2411     for antineutrino propagation. Thus, a model of the Earth's density is required  
 2412     for neutrino propagation. Following the official SK-only methodology [206],  
 2413     this analysis uses the Preliminary Reference Earth Model (PREM) [207] which  
 2414     provides piecewise cubic polynomials as a function of the Earth's radius. This  
 2415     density profile is illustrated in Figure 7.1. As the propagator requires layers of  
 2416     constant density, the SK methodology approximates the PREM model by using  
 2417     four layers of constant density [206], detailed in Table 7.1.

Layer	Outer Radius [km]	Density [ $\text{g}/\text{cm}^3$ ]	Chemical composition (Z/A)
Inner Core	1220	13	$0.468 \pm 0.029$
Outer Core	3480	11.3	$0.468 \pm 0.029$
Lower Mantle	5701	5.0	0.496
Transition Zone	6371	3.3	0.496

**Table 7.1:** Description of the four layers of the Earth invoked within the constant density approximation of the PREM model [207].

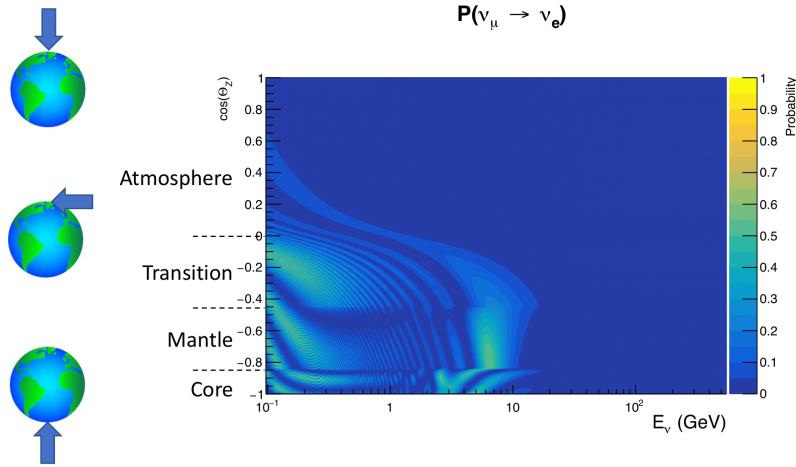
2418     The atmospheric neutrino oscillation probabilities can be presented as two di-  
 2419     mensional “oscillograms” as illustrated in Figure 7.2. The distinct discontinuities,  
 2420     as a function of  $\cos(\theta_Z)$ , are due to the discontinuous density in the PREM model.



**Figure 7.1:** The density of the Earth given as a function of the radius, as given by the PREM model (Black), and the constant density four-layer approximation (Blue), as used in the official SK-only analysis.

Atmospheric neutrinos have sensitivity to  $\delta_{CP}$  through the overall event rate. Figure 7.3 illustrates the difference in oscillation probability between CP-conserving ( $\delta_{CP} = 0.$ ) and a CP-violating ( $\delta_{CP} = -1.601$ ) value taken from Asimov A oscillation parameter set (Table 2.2). The result is a complicated oscillation pattern in the appearance probability for sub-GeV upgoing neutrinos. The detector does not have sufficient resolution to resolve these individual patterns so the sensitivity to  $\delta_{CP}$  for atmospheric neutrinos comes via the overall normalisation of these events.

The presence of matter means that the effect  $\delta_{CP}$  has on the oscillation probability is not equal between neutrinos and antineutrinos. Furthermore, the interaction cross-section for neutrinos is larger than for antineutrinos so the two effects have to be disentangled. These effects are further convoluted by detector efficiencies as SK cannot distinguish neutrinos and antineutrinos well. Furthermore, the sample selections discussed in section 6.1 have different efficiencies for neutrino-antineutrino selections. All of these effects lead to a difference in the number of neutrinos detected compared to antineutrinos. This changes how the  $\delta_{CP}$



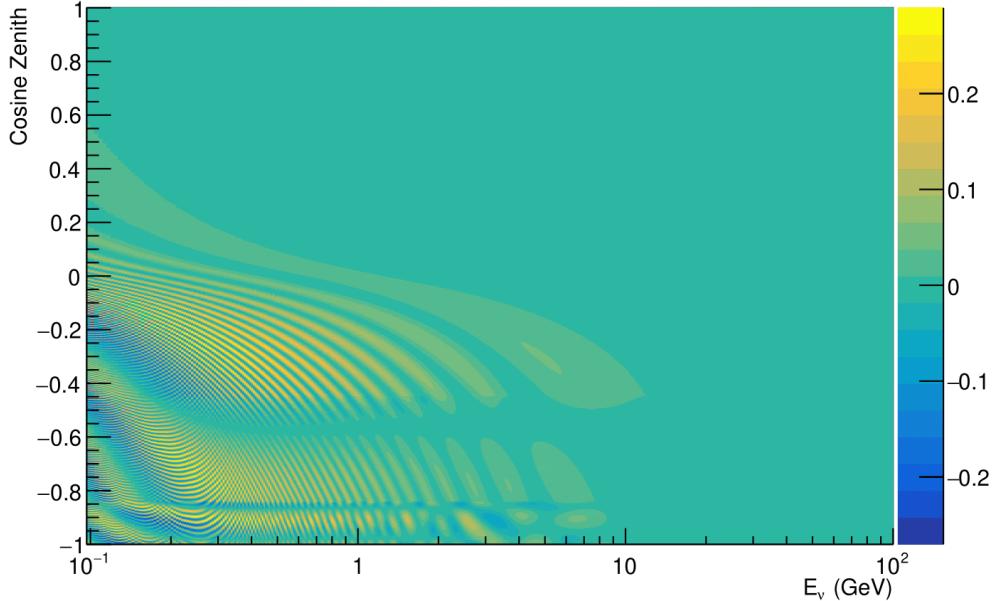
**Figure 7.2:** An “oscillogram” that depicts the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability as a function of neutrino energy and cosine of the zenith angle. The zenith angle is defined such that  $\cos(\theta_Z) = 1.0$  represents neutrinos that travel from directly above the detector. The four-layer constant density PREM model approximation is used and Asimov A oscillation parameters are assumed (Table 2.2).

normalisation term is observed, resulting in a very complex sensitivity to  $\delta_{CP}$ .

The vacuum and matter oscillation probabilities for  $P(\nu_e \rightarrow \nu_e)$  and  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  are presented in Figure 7.4, where the PREM model has been assumed. The oscillation probability for both neutrinos and antineutrinos is affected in the presence of matter. However, the resonance effects around  $O(5)\text{GeV}$  only occur for neutrinos in the normal mass hierarchy and antineutrinos in the inverse mass hierarchy. The exact position and amplitude of the resonance depend on  $\sin^2(\theta_{23})$ , further increasing the atmospheric neutrinos’ sensitivity to the parameter.

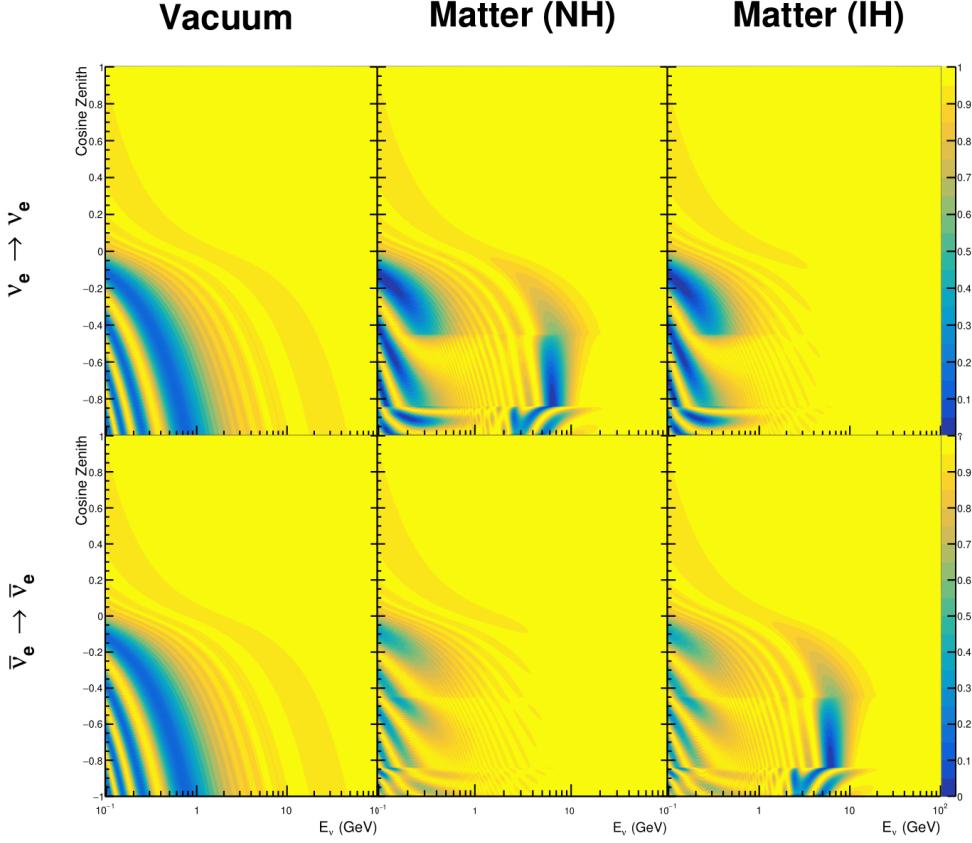
As the T2K beam flux is centered at the first oscillation maximum ( $E_\nu = 0.6\text{GeV}$ ), the sensitivity to  $\delta_{CP}$  is predominantly observed as a change in the event-rate of e-like samples in  $\nu/\bar{\nu}$  modes. Figure 7.5 illustrates the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability for a range of  $\delta_{CP}$  values. A circular modulation of the first oscillation peak (in both magnitude and position) is observed when varying throughout the allowable values of  $\delta_{CP}$ . The CP-conserving values of  $\delta_{CP} = 0, \pi$  have a lower(higher) oscillation maximum than the CP-violating values of  $\delta_{CP} = -\pi/2 (\delta_{CP} = \pi/2)$ . A sub-dominant shift in the energy of the oscillation peak is

$$\mathbf{P}(\nu_\mu \rightarrow \nu_e; \delta_{CP} = -1.601) - \mathbf{P}(\nu_\mu \rightarrow \nu_e; \delta_{CP} = 0.)$$



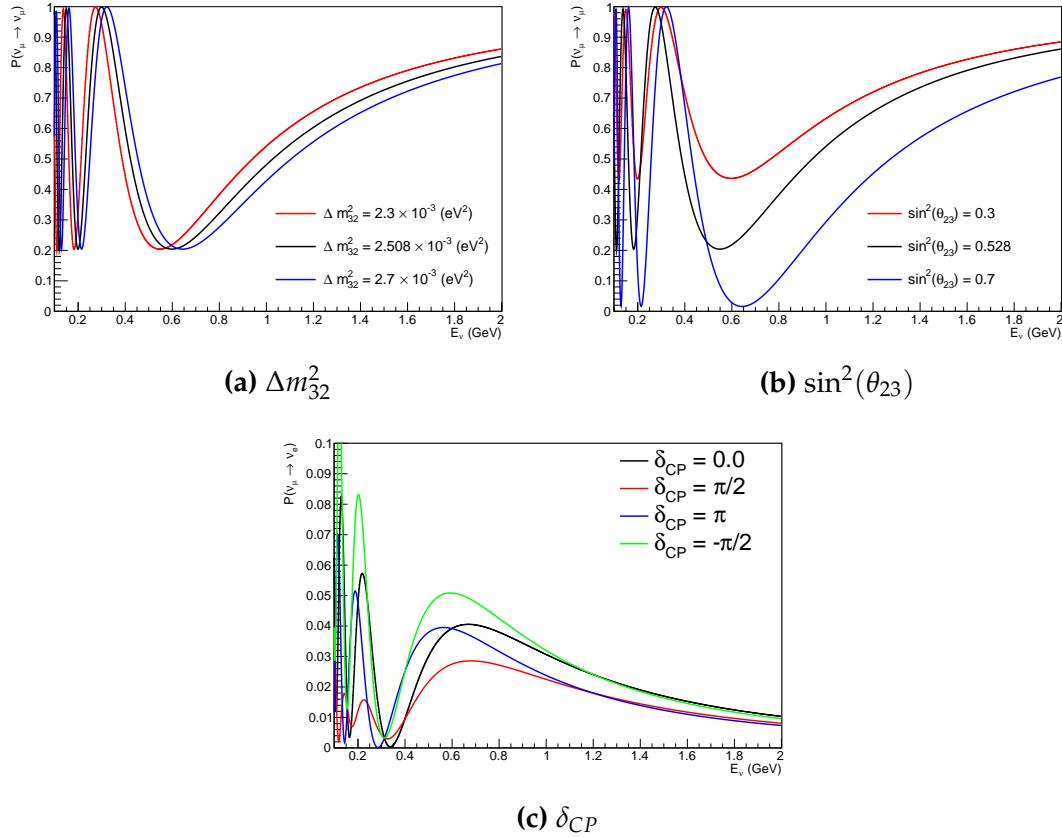
**Figure 7.3:** The effect of  $\delta_{CP}$  for atmospheric neutrinos given in terms of the neutrino energy and zenith angle. This oscillogram compares the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability for a CP conserving ( $\delta_{CP} = 0.0$ ) and a CP violating ( $\delta_{CP} = -1.601$ ) value taken from the Asimov A parameter set. The other oscillation parameters assume the Asimov A oscillation parameter set given in Table 2.2.

2453 also present, which aids in separating the two CP-conserving values of  $\delta_{CP}$ .  
 2454 T2K's sensitivity to  $\sin^2(\theta_{23})$  and  $\Delta m_{32}^2$  is observed as a shape-based variation  
 2455 of the muon-like samples, as illustrated in Figure 7.5. The value of  $\Delta m_{32}^2$  laterally  
 2456 shifts the position of the oscillation dip (around  $E_\nu \sim 0.6\text{GeV}$ ) in the  $P(\nu_\mu \rightarrow \nu_\mu)$   
 2457 oscillation probability. A variation of  $\sin^2(\theta_{23})$  is predominantly observed as  
 2458 a vertical shift of the oscillation dip with second-order horizontal shifts being  
 2459 due to matter effects. The beam neutrinos have limited sensitivity to matter  
 2460 effects due to the relatively shorter baseline as well as the Earth's mantle being  
 2461 a relatively low-density material (as compared to the Earth's core). For some  
 2462 values of  $\delta_{CP}$ , the degeneracy in the number of e-like events allows the mass  
 2463 hierarchy to be broken. This leads to a  $\delta_{CP}$ -dependent mass hierarchy sensitivity  
 2464 which can be seen in Figure 7.6.



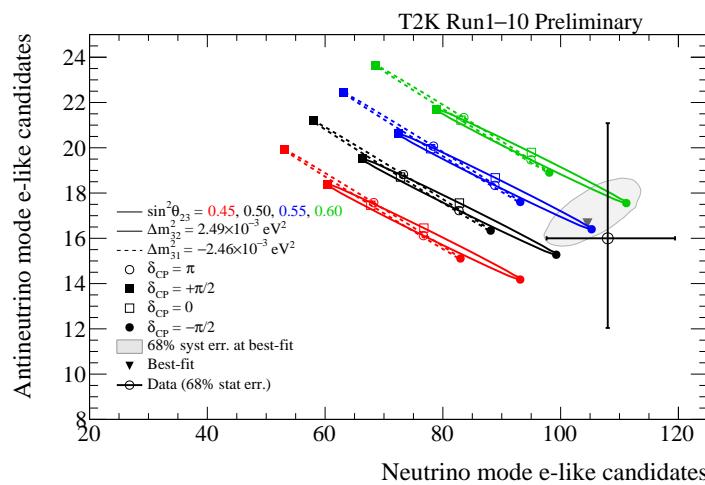
**Figure 7.4:** An illustration of the matter-induced effects on the oscillation probability, given as a function of neutrino energy and zenith angle. The top row of panels gives the  $P(\nu_e \rightarrow \nu_e)$  oscillation probability and the bottom row illustrates the  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  oscillation probability. The left column highlights the oscillation probability in a vacuum, whereas the middle and right column represents the oscillation probabilities when the four-layer fixed density PREM model is assumed. All oscillation probabilities assume the “Asimov A” set given in Table 2.2, but importantly, the right column sets an inverted mass hierarchy. The “matter resonance” effects at  $E_\nu \sim 5\text{GeV}$  can be seen in the  $P(\nu_e \rightarrow \nu_e)$  for normal mass hierarchy and  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  for inverted hierarchy.

Whilst all oscillation channels should be included for completeness, the computational resources required to run a fit are limited and any reasonable approximations which reduce the number of oscillation probability calculations that need to be made should be applied. The  $\nu_e \rightarrow \nu_{e,\mu,\tau}$  (and antineutrino equivalent) oscillations can be ignored for beam neutrinos as the  $\nu_e/\bar{\nu}_e$  fluxes are approximately two orders of magnitude smaller than the corresponding  $\nu_\mu/\bar{\nu}_\mu$  flux. Furthermore, as the peak neutrino energy of the beam is well below the



**Figure 7.5:** The oscillation probability for beam neutrino events given as a function of neutrino energy. All oscillation parameters assume the “Asimov A” set given in Table 2.2 unless otherwise stated. Each panel represents a change in one of the oscillation parameters whilst keeping the remaining parameters fixed.

threshold for charged current tau production ( $E_\nu = 3.5\text{GeV}$  [58]), only a small proportion of the neutrinos produced in the beam have the required energy. For the few neutrinos that have sufficient energy, the oscillation probability is very small due to their energy being well above the oscillation maximum (small value of  $L/E$ ). Whilst these approximations have been made for the beam neutrinos, the atmospheric flux of  $\nu_e$  is of the same order of magnitude as the  $\nu_\mu$  flux and the energy distribution of atmospheric neutrinos extends well above the tau production threshold. These events can have non-negligible oscillation probabilities due to the further distance they travel.



**Figure 7.6:** The number of electron-like events in the FHC and RHC operating mode of the beam, as a function of the oscillation probabilities. Both normal hierarchy (Solid) and inverse hierarchy (Dashed) values of  $\Delta m_{32}^2$  are given.

## 2481 7.2 Treatment of Fast Oscillations

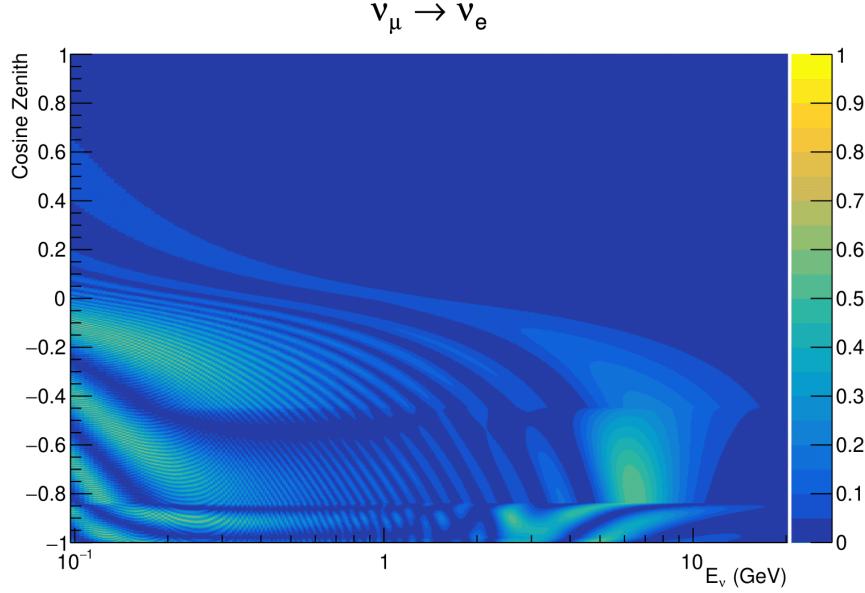
2482 As shown in Figure 7.7, atmospheric neutrino oscillations have a significantly  
 2483 more complex structure for upgoing neutrinos with energy below 1GeV. This is  
 2484 because the  $L/E$  dependence of the oscillation probability in this region induces  
 2485 rapid variations for small changes in  $L$  or  $E$ . As discussed in section 7.1, this is also  
 2486 the region in which atmospheric neutrinos have sensitivity to  $\delta_{CP}$ . In practice,  
 2487 the direction of the neutrino is inferred from the direction of the final state  
 2488 particles traveling in the detector. The correlation between these two directions  
 2489 can be particularly weak for low-energy neutrino interactions. This creates a  
 2490 distinct difference from the beam neutrinos where the position of the source  
 2491 is very precisely known.

2492 As a consequence of the unresolvable structure, an event rate consistent with  
 2493 the averaged oscillation probability is observed in the subGeV upgoing region.  
 2494 This creates a computational problem: A significantly large amount of Monte  
 2495 Carlo statistics would be required to accurately predict the number of events  
 2496 if Monte Carlo averaging was the only technique used. This section describes  
 2497 the ‘sub-sampling’ approach developed for this analysis and compares it to the  
 2498 methodology used within the SK-only analysis.

2499 **DB: Figure out what this actually means** The official SK-only analysis uses the  
 2500 osc3++ oscillation parameter fitter [206]. To perform the fast oscillation averaging,  
 2501 it uses a ‘nearest-neighbour’ technique. For a given Monte Carlo neutrino event,  
 2502 the nearest twenty Monte Carlo neighbours in reconstructed lepton momentum  
 2503 and zenith angle are found and a distribution of their neutrino energies is built.  
 2504 The RMS,  $\sigma$ , of this distribution is then used to compute an average oscillation  
 2505 probability for the given neutrino Monte Carlo event.

2506 For the  $i^{th}$  event, the oscillation weight is calculated as

$$W_i = \frac{1}{5}P(E_i, \bar{L}_i) + \frac{1}{5} \sum_{\beta=-1, -0.5, 0.5, 1} P(E_i + \beta\sigma_i, L_\beta), \quad (7.3)$$



**Figure 7.7:** The oscillation probability  $P(\nu_\mu \rightarrow \nu_e)$ , given as a function of neutrino energy and zenith angle, which highlights an example of the “fast” oscillations in the sub-GeV upgoing region.

2507 where  $P(E, L)$  is the oscillation probability calculation for neutrino energy  
 2508  $E$  and path length  $L$  and the two path lengths,  $\bar{L}_i$  and  $L_\beta$  are discussed below.  
 2509 All of the oscillation probability calculations are performed with a fixed zenith  
 2510 angle such that the same density profile is used.

2511 The uncertainty in the production height is controlled by using an “average”  
 2512 production height,  $\bar{L}_i$ , which represents the average path length computed using  
 2513 twenty production heights taken from the Honda flux model’s prediction [52].  
 2514 For a given event, the production heights are sampled in steps of 5% of their  
 2515 cumulative distribution function.  $L_\beta$  values are similarly calculated but instead  
 2516 use different combinations of four production heights,

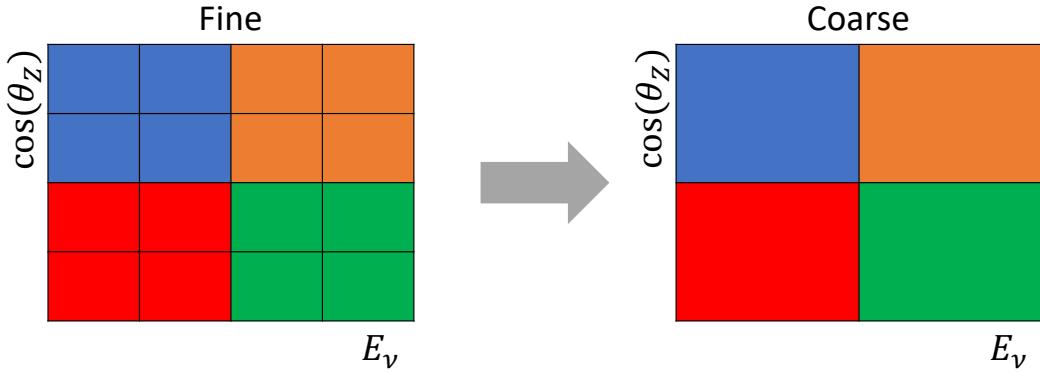
$$\begin{aligned}
 L_{-1.0} &= \frac{1}{4}L(45, 50, 55, 60), \\
 L_{-0.5} &= \frac{1}{4}L(35, 40, 65, 70), \\
 L_{+0.5} &= \frac{1}{4}L(25, 30, 75, 68), \\
 L_{+1.0} &= \frac{1}{4}L(15, 20, 85, 89).
 \end{aligned} \tag{7.4}$$

This averaging technique works because of the inference between the zenith angle and the reconstructed direction of final state particles in the detector. For low-energy neutrinos, where the resolution of the true neutrino direction is poor,  $\sigma_i$  will be large, resulting in significant averaging effects. Contrary to this, the inferred direction of high-energy neutrinos will be much closer to the true value, meaning that  $\sigma_i$  will be smaller, culminating in small averaging effects.

In practice, these calculations are performed prior to the fit as only oscillation parameters at fixed points are considered. The MCMC technique used in this thesis requires oscillation probabilities to be evaluated at arbitrary parameter values, not known *a priori*. Calculating the five oscillation probabilities per event required by the SK technique is computationally infeasible, so a different averaging technique is used. However, the concept of the averaging technique can be taken from it.

To perform a similar averaging as the SK analysis, a sub-sampling approach using binned oscillograms has been devised. A coarsely binned oscillogram is defined in  $\cos(\theta_Z)$  and  $E_\nu$ . For a given set of oscillation parameters, a single oscillation probability will be assigned to each coarse bin. This value will then apply to all Monte Carlo events which fall into that bin. To assign these oscillation probabilities, the probability is calculated at  $N \times N$  points on a grid within a particular bin. This ensemble of oscillation probabilities is averaged to define the coarse bin's oscillation probability, assuming a flat prior in  $E_\nu$  and  $\cos(\theta_Z)$  within the bin. Figure 7.8 illustrates the  $N = 2$  example where the assigned value to a coarse bin is the average of the four fine bins which fall in that coarse bin. Whilst the coarse bin edges do not have to be linear on either axis, the sub-division of the fine bins is linear within the range of a coarse bin.

The coarse binning is defined with  $67 \times 52$  bins in true neutrino energy  $\times$  cosine zenith. It is picked to be identical to that provided in [208]. In general, the binning is logarithmically spaced in neutrino energy but has some hand-picked bin edges around the matter resonance to smoothly increased the bin density. This is to avoid smearing this region which can be well sampled by the Monte

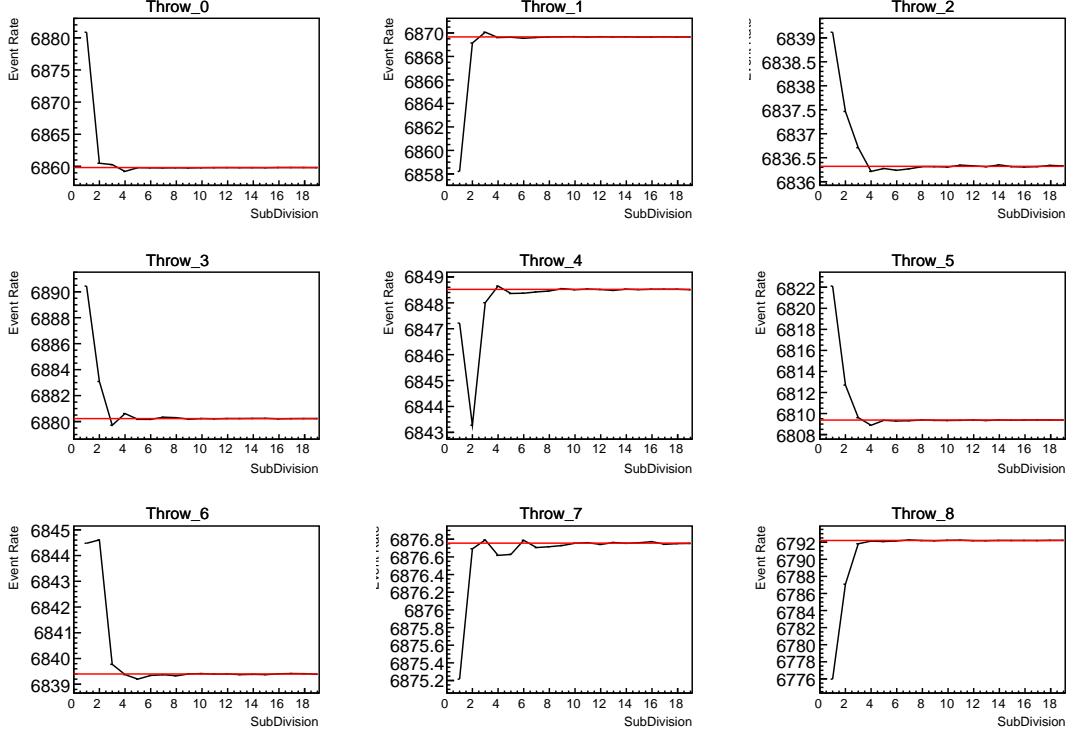


**Figure 7.8:** Illustration of the averaging procedure for  $N = 2$ . The oscillation probabilities calculated on the finer left binning are averaged to obtain the oscillation probabilities in the coarser right binning. These averaged oscillation probabilities with the coarser binning are then applied to each event during the fit.

2547 Carlo. The cosine zenith binning is approximately linearly spaced across the  
 2548 allowable range but the values of layer transitions are hit precisely:  $-0.8376$  (core-  
 2549 mantle) and  $-0.4464$  (mantle/transition zone). Bins are spread further apart for  
 2550 downgoing events as this is a region unaffected by the fast oscillation wavelengths  
 2551 and reduces the total number of calculations required to perform the calculation.

2552 The choice of  $N$  is justified based on two studies. Firstly, the variation of event  
 2553 rates of each sample is studied as a function of  $N$ . For a given set of oscillation  
 2554 parameters thrown from the PDG prior constraints (detailed in Table 2.1), the  
 2555 oscillation probabilities are calculated using a given value of  $N$ . Each sample  
 2556 is re-weighted and the event rate is stored. The value of  $N$  is scanned from  
 2557 1, which corresponds to no averaging, to 19, which corresponds to the largest  
 2558 computationally viable subdivision binning. The event rate of each sample at  
 2559 large  $N$  is expected to converge to a stationary value due to the fine binning fully  
 2560 sampling the small-scale structure. Figure 7.9 illustrates this behaviour for the  
 2561 SubGeV\_elike\_0dcy sample for 9 different throws of the oscillation parameters.

2562 Denoting the event rate for one sample for a given throw  $t$  at each  $N$  by  $\lambda_t^N$ ,  
 2563 the average over all considered  $N$  values ( $\bar{\lambda}_t = \frac{1}{24} \sum_{N=1}^{24} \lambda_t^N$ ) is computed. The



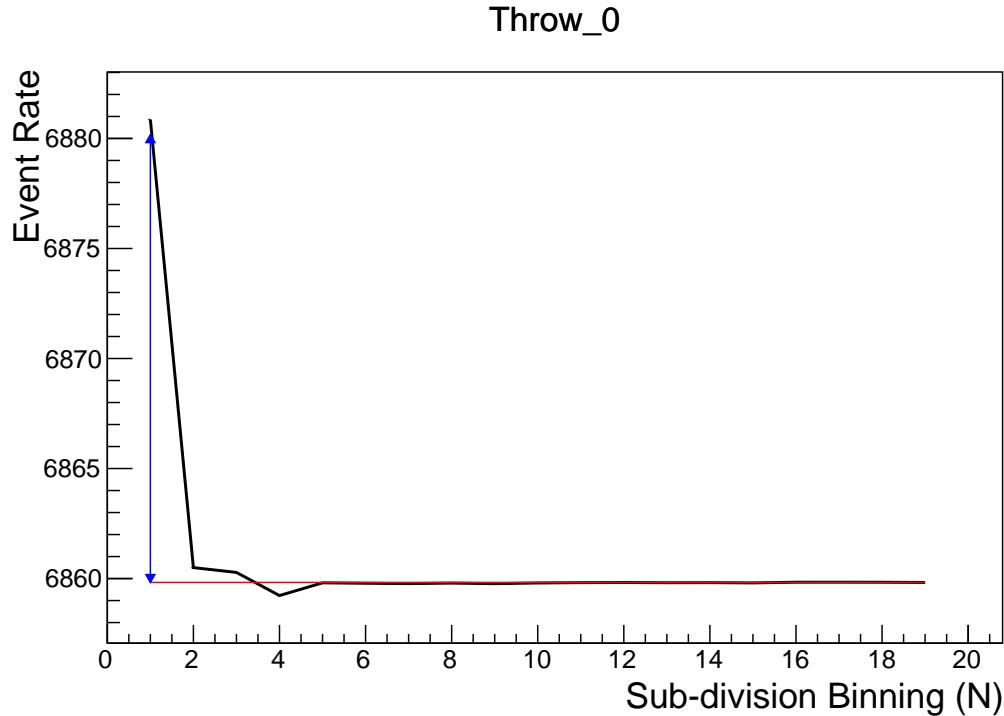
**Figure 7.9:** Event rate of the SubGeV\_elike\_0dcy sample as a function of the number of sub-divisions per coarse bin. Each subplot represents the event rate of the sample at a different oscillation parameter set thrown from the PDG priors detailed in Table 2.1. The red line in each subplot represents the mean of the event rate over the different values of sub-divisions for that particular oscillation parameter throw.

variance in the event rate at each  $N$  is then calculated as

$$\text{Var}[\lambda^N] = \frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t)^2 - \left[ \frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t) \right]^2. \quad (7.5)$$

In practice, the following procedure is undertaken. For a particular throw, the difference between the event rate at a particular choice of  $N$  and the mean of the distribution is calculated. This is illustrated in Figure 7.10. This value is then calculated for all the 2000 throws, generating a distribution of  $\lambda_t^N - \bar{\lambda}_t$ . This is repeated for each of the values of  $N$  considered within this study. The distributions of this value, for  $N = \{1, 5\}$ , are given in Figure 7.11. As expected, the distribution gets narrower and tends towards zero for the higher values of  $N$ .

The aim of the study is to find the lowest value of  $N$  such that this variance is below 0.001. This utilises the width of the distributions given in Figure 7.11.

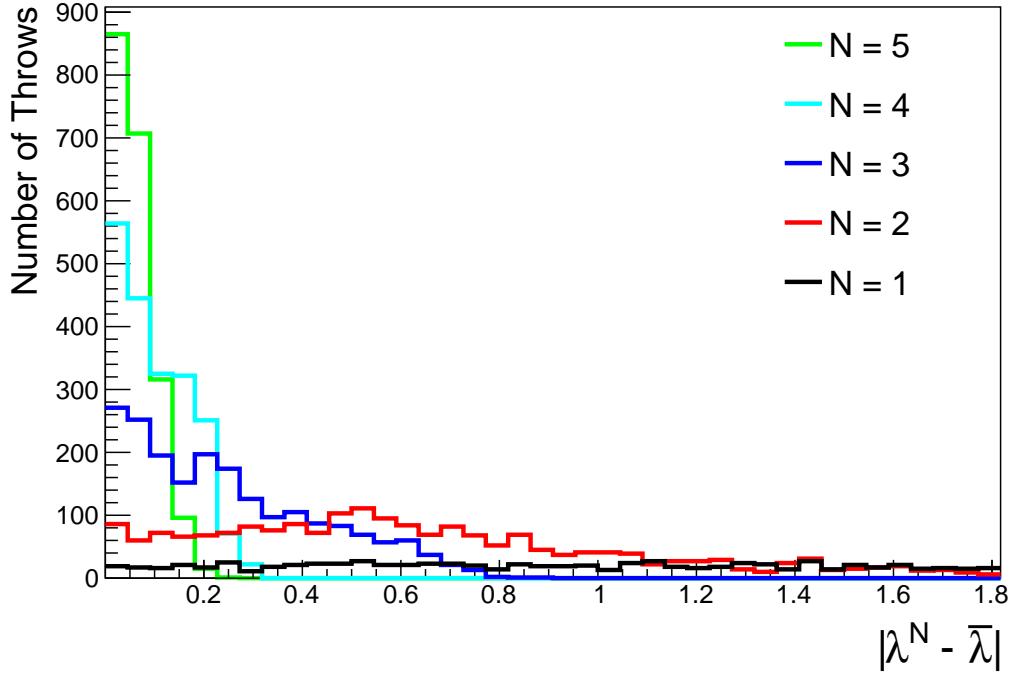


**Figure 7.10:** Event rate of the SubGeV\_elike\_0dcy sample, for a particular oscillation parameter throw, as a function of the number of sub-divisions,  $N$ , per coarse bin. The difference between the mean event rate (red),  $\bar{\lambda}$ , and the event rate at  $N = 1$ ,  $\lambda^{N=1}$  is defined as  $\lambda^N - \bar{\lambda}$  and illustrated by the blue arrow.

2574 This is the typical threshold used by T2K fitters to validate systematic imple-  
 2575 mentation so has been set as the same criteria. The results of this study for  
 2576 each atmospheric sample used within this thesis are illustrated in Figure 7.12 for  
 2577 2000 throws of the oscillation parameters. As can be seen, the variance is below  
 2578 the threshold at  $N = 10$ , and is driven primarily by the SubGeV\_mulike\_1dcy  
 2579 and SubGeV\_elike\_0dcy samples.

2580 The second study to determine the value of  $N$  is as follows. The likelihood  
 2581 for each sample is computed against an Asimov data set created with Asimov  
 2582 A oscillation parameters (Table 2.2). Following Equation 7.5, the variance of  
 2583 the log-likelihood over all considered  $N$  is computed. The results are shown  
 2584 in Figure 7.13.

2585 A choice of  $N = 10$  sub-divisions per coarse bin has a variance in both  
 2586 event rate and log-likelihood residuals less than the required threshold of 0.001.



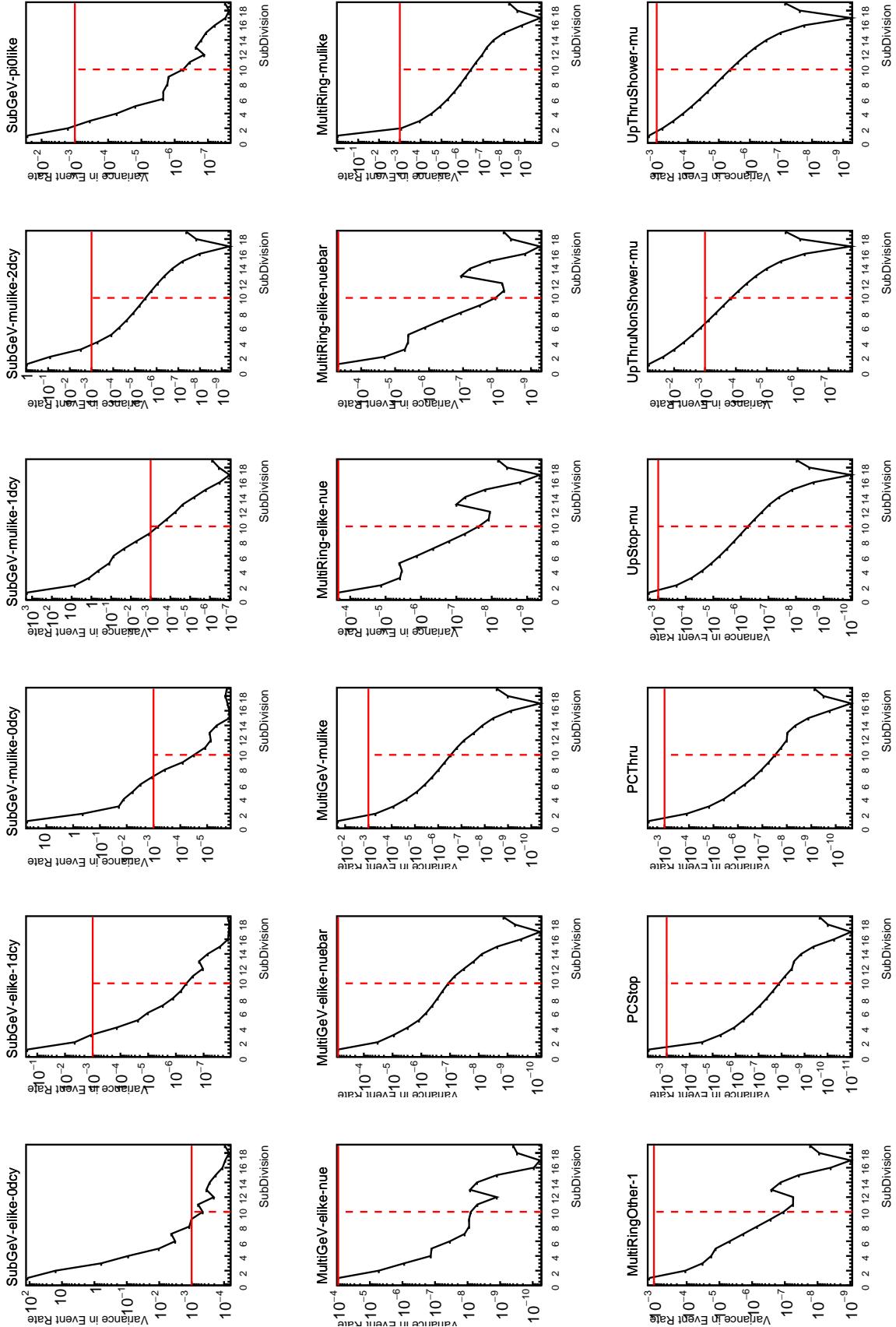
**Figure 7.11:** The distribution of  $\lambda^N - \bar{\lambda}$  for various values of  $N$ . As expected, the distribution gets narrower for larger values of  $N$ .

2587 The largest value of the likelihood variance is of order  $10^{-7}$ , corresponding to  
 2588 an error on the log-likelihood of about  $3 \times 10^{-4}$  which is small enough to be  
 2589 negligible for the oscillation analysis.

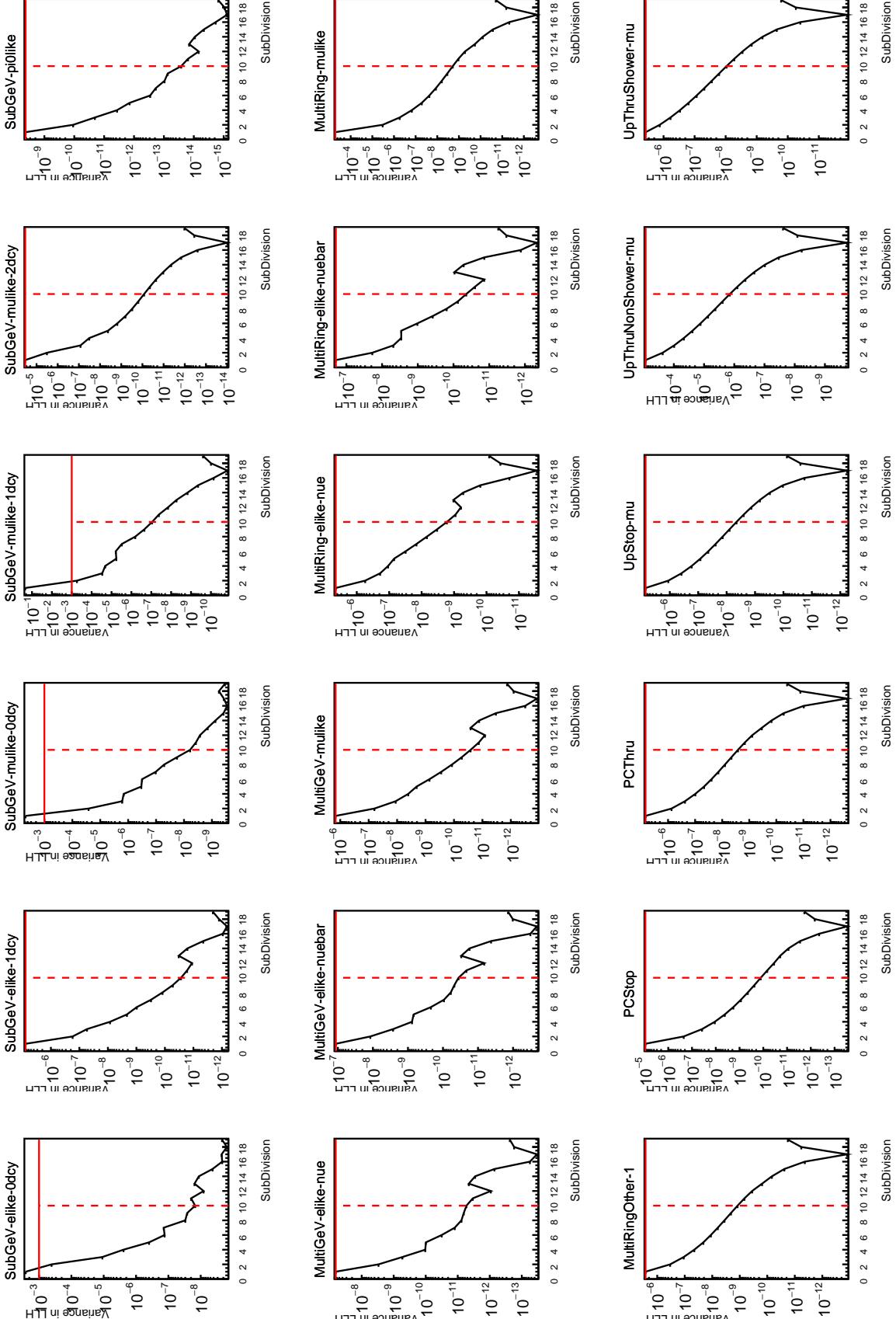
2590 Figure 7.14 illustrates the effect of the smearing using  $N = 10$ . The fast oscil-  
 2591 lations in the sub-GeV upgoing region have been replaced with a normalisation  
 2592 effect whilst the large matter resonance structure remains.

### 2593 7.3 Calculation Engine

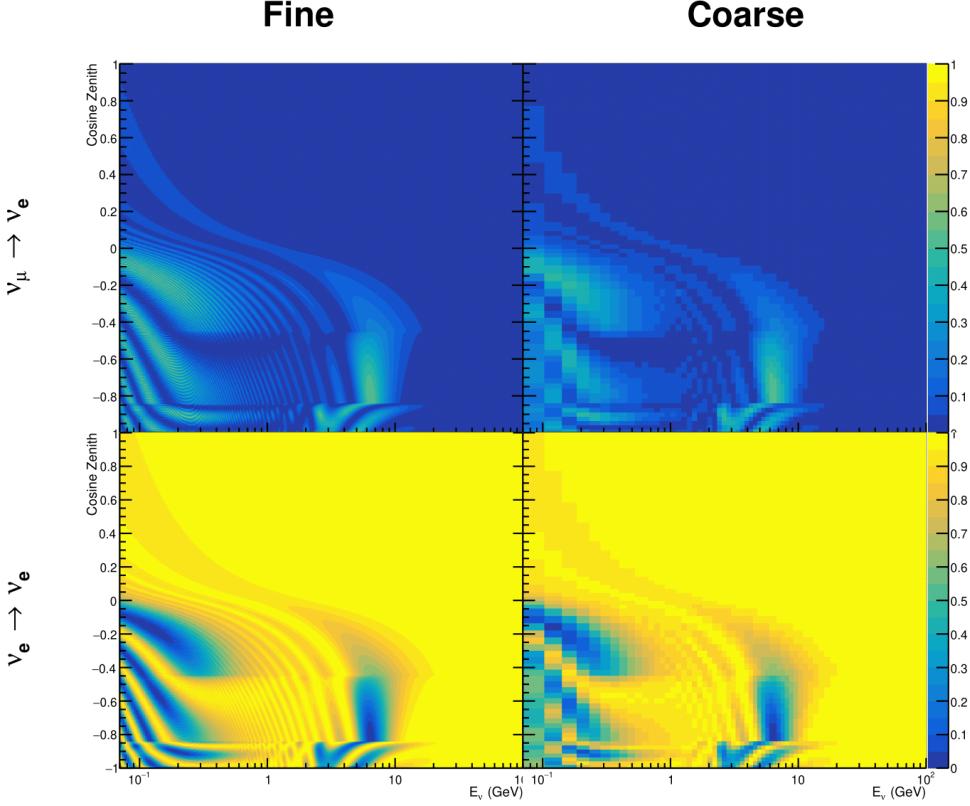
2594 As previously discussed in section 7.2, the calculation of oscillation probabilities  
 2595 is performed at run-time. Consequently, the time per calculation is crucial for fit  
 2596 performance. The initial fitting framework used for this analysis was developed  
 2597 with ProbGPU [209]. This is a GPU-only implementation of the prob3 engine  
 2598 [210]. It is primarily designed for neutrino propagation in a beam experiment



**Figure 7.12:** Variance of event rate for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line indicates the variance at a sub-division  $N = 10$ .



**Figure 7.13:** Variance of sample likelihood, when compared to ‘Asimov data’ set at Asimov A, for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line is a graphical indication of the variance at a sub-division  $N = 10$ .

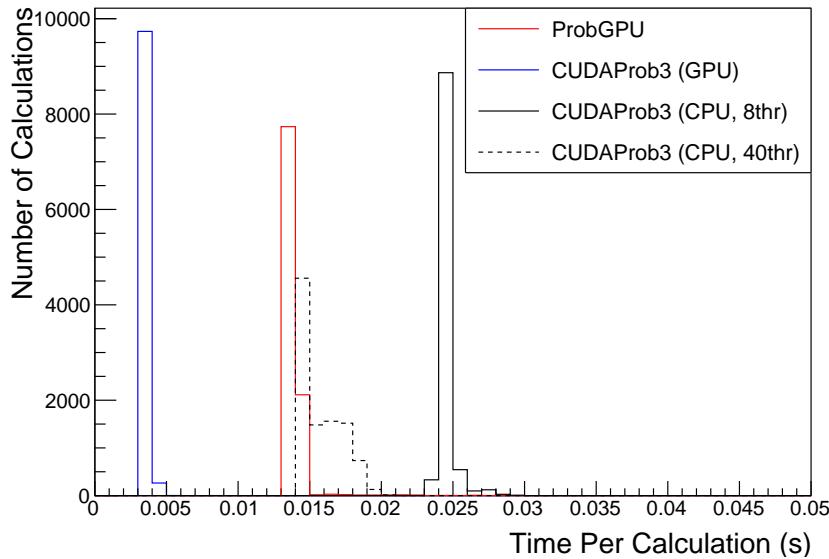


**Figure 7.14:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives the “fine” binning used to calculate the oscillation probabilities and the right column illustrates the “coarse” binning used to reweight the Monte Carlo events. The fine binning choice is given with  $N = 10$ , which was determined to be below the threshold from Figure 7.12 and Figure 7.13.

2599 (single layer of constant density) with the atmospheric propagation code not  
 2600 being used prior to the analysis in this thesis.

2601 Another engine, CUDAProb3 [211], has been interfaced with the fitting frame-  
 2602 work used in this analysis. This interfacing was done by the author of this  
 2603 thesis. It has been specifically optimised for atmospheric neutrino oscillation  
 2604 calculation so does not contain the code to replace the beam oscillation calculation.  
 2605 The engine utilises object-orientated techniques as compared to the functional  
 2606 implementation of ProbGPU. This allows the energy and cosine zenith arrays to  
 2607 be kept on GPU memory, rather than having to load these arrays onto GPU

2608 memory for each calculation. Reducing the memory transfer between CPU and  
 2609 GPU significantly reduces the time required for calculation. This can be seen  
 2610 in Figure 7.15, where the GPU implementation of CUDAProb3 is approximately  
 2611 three times faster than the ProbGPU engine.



**Figure 7.15:** The calculation time taken to both calculate the oscillation probabilities and fill the “coarse” oscillograms, following the technique given in section 7.2, for the CUDAProb3 and ProbGPU (Red) calculation engines. CUDAProb3 has both a GPU (Blue) and CPU (Black) implementation, where the CPU implementation is multi-threaded. Therefore, 8-threads (solid) and 40-threads (dashed) configurations have been tested. Prob3, which is a CPU single-thread implementation has a mean step time of 1.142s.

2612 Another significant advantage of CUDAProb3 is that it contains a CPU mul-  
 2613 tithreaded implementation which is not possible with the ProbGPU or prob3  
 2614 engines. This eliminates the requirement for GPU resources when submitting  
 2615 jobs to batch systems. As illustrated in Figure 7.15, the calculation speed depends  
 2616 on the number of available threads. Using 8 threads (which is typical of the  
 2617 batch systems being used) is approximately twice as slow as the ProbGPU engine  
 2618 implementation, but would allow the fitting framework to be run on many more  
 2619 resources. This fact is utilised for any SK-only fits but GPU resources are required  
 2620 for any fits which include beam samples due to the ProbGPU requirement. Based  
 2621 on the benefits shown by the implementation in this section, efforts are being

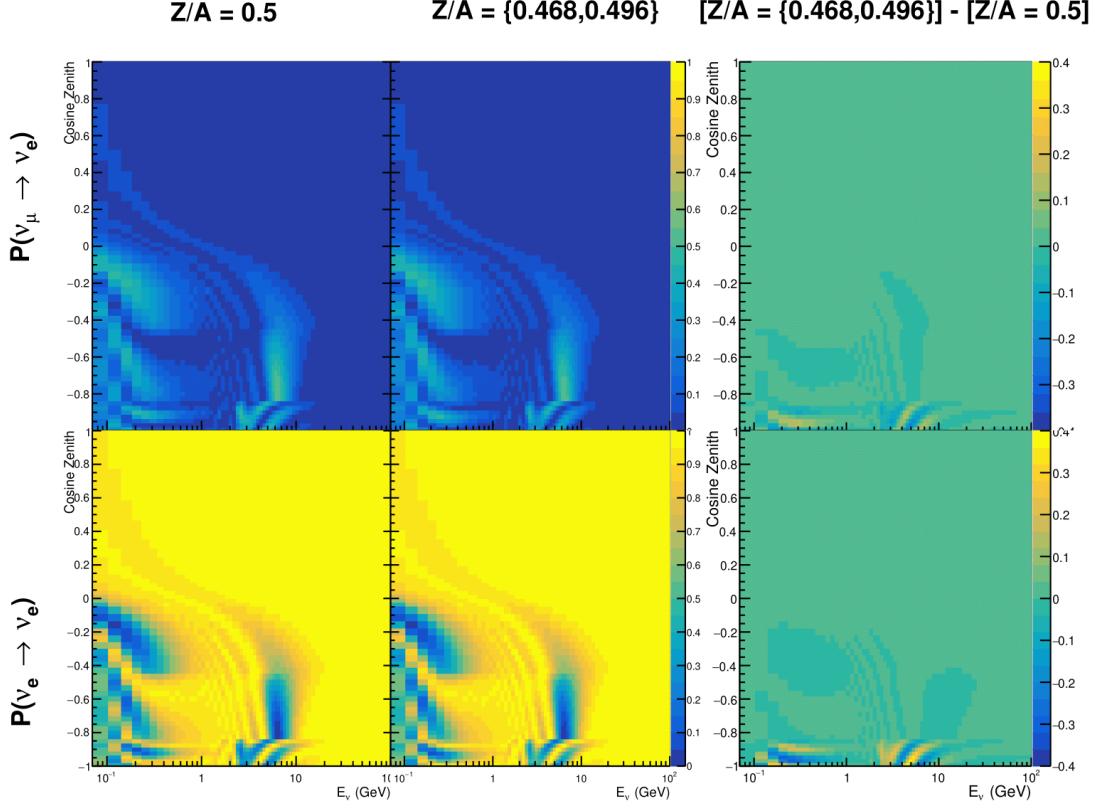
2622 placed into including linear propagation for beam neutrino propagation into  
2623 the engine [212].

## 2624 7.4 Matter Density Profile

2625 For an experiment observing neutrinos propagating through the Earth, a model  
2626 of the Earth's density profile is required. The model used within this analysis is  
2627 based on the Preliminary Reference Earth Model (PREM) [207], as illustrated in  
2628 Figure 7.1. Table 7.1 documents the density and radii of the layers used within  
2629 the constant density approximation used by the SK-only analysis [206]. The  
2630 density measurements provided in the PREM model are provided in terms of  
2631 mass density, whereas neutrino oscillations are sensitive to the electron number  
2632 density. This value can be computed as the product of the chemical composition,  
2633 or the  $Z/A$  value, and the mass density of each layer. Currently, the only way  
2634 to measure the chemical composition value for layers close to the Earth's core  
2635 is through neutrino oscillations. The chemical composition of the upper layers  
2636 of the Earth's Mantle and the Transition zone is well known due to it being  
2637 predominantly pyrolite which has a chemical composition value of 0.496 [213].  
2638 The chemical composition dial for the core layers is set to a value of 0.468, as  
2639 calculated in [214]. As this value is less well known, it is assigned a Gaussian error  
2640 with a standard deviation equivalent to the difference in chemical composition  
2641 in core and mantle layers. Figure 7.16 illustrates the effect of moving from  
2642 the  $Z/A = 0.5$  method which is used in the official SK-only analysis to these  
2643 more precise values.

2644 The beam oscillation probability in this thesis uses a baseline of 295km, density  
2645  $2.6\text{g}/\text{cm}^3$ , and chemical composition 0.5 as is done by the official T2K-only  
2646 analysis [215].

2647 For a neutrino with given  $E_\nu$ ,  $\cos(\theta_Z)$ , the oscillation probability calculation  
2648 engine must be passed a list of the matter regions that the neutrino traversed,  
2649 with the path length and fixed density in each region. However, a neutrino  
2650 passing through the earth experiences a range of radii, and thus a range of



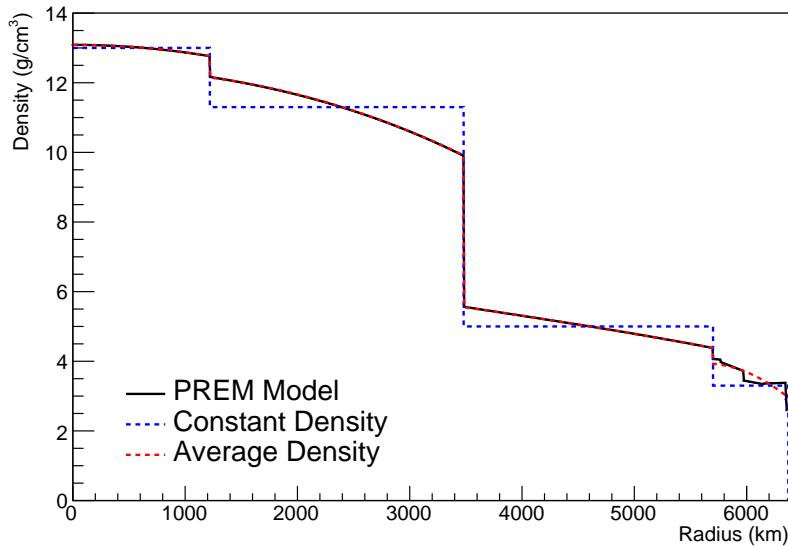
**Figure 7.16:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the constant  $Z/A = 0.5$  approximation which is used in the official SK-only analysis. The middle column gives the probabilities where  $Z/A = [0.468, 0.498]$  values are used, as given in Table 7.1. The right column illustrates the difference in oscillation probability between the two different techniques.

2651 densities, in each region. In the SK-only analysis, the earth density model used  
 2652 is piecewise-constant, thereby ignoring this effect. For this thesis, the density  
 2653 values for the calculation engine are found by averaging the earth density along  
 2654 the neutrino's path in each layer,

$$\langle \rho \rangle_i = \frac{1}{t_{i+1} - t_i} \int_{t_i}^{t_{i+1}} \rho(t) dt \quad (7.6)$$

2655 where  $t_i$  are the intersection points between each layer and  $t$  is the path length  
 2656 of the trajectory across the layer. This leads to an improved approximation.  
 2657 For this averaging, the simplification of the PREM model developed in [216] is

used. The layers of the prem model are combined into four to reduce calculation time, with a quadratic fit to each section. This fit was not performed by the author of the thesis and is documented in [208]. The coefficients of the quadratic fit to each layer are given in Table 7.2 with the final distribution illustrated in Figure 7.17. The quadratic approximation is clearly much closer to the PREM model as compared to the constant density approximation.



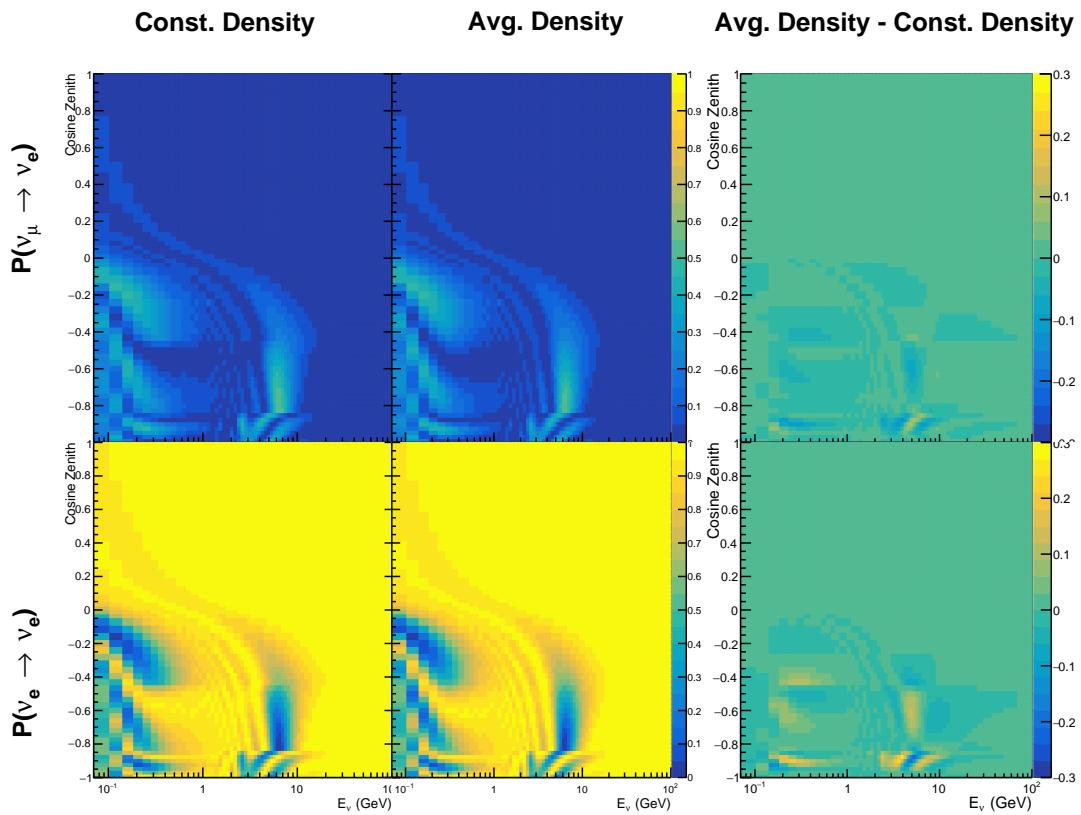
**Figure 7.17:** The density of the Earth given as a function of the radius, as given by the PREM model (Black), the constant density four-layer approximation (Blue), as used in the official SK-only analysis, and the quadratic approximation of the PREM model (Red).

Layer	Outer Radius [km]	Density [g/cm <sup>3</sup> ]
Inner Core	1220	$13.09 - 8.84x^2$
Outer Core	3480	$12.31 + 1.09x - 10.02x^2$
Lower Mantle	5701	$6.78 - 1.56x - 1.25x^2$
Transition Zone	6371	$-50.42 + 123.33x - 69.95x^2$

**Table 7.2:** The quadratic polynomial fits to the PREM model for four assumed layers of the PREM model. The fit to calculate the coefficients is given in [208], where  $x = R/R_{Earth}$ .

The effect of using the quadratic density per  $\cos(\theta_Z)$  model is highlighted in Figure 7.18. The slight discontinuity in the oscillation probability around  $\cos(\theta_Z) \sim -0.45$  in the fixed density model, which is due to the transition to mantle layer boundary, has been reduced. This is expected as the difference in

the density across this boundary is significantly smaller in the quadratic density model as compared to the constant density model. Whilst the difference in density across the other layer transitions is reduced, there is still a significant difference. This means the discontinuities in the oscillation probabilities remain but are significantly reduced. However, as the quadratic density approximation matches the PREM model well in this region, these discontinuities are due to the Earth model rather than an artifact of the oscillation calculation.



**Figure 7.18:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the four-layer constant density approximation is used. The middle column gives the probabilities where the density is integrated over the trajectory, using the quadratic PREM approximation, for each  $\cos(\theta_Z)$  is used. The right column illustrates the difference in oscillation probability between the two different techniques.

## 2675 7.5 Production Height Averaging

2676 As discussed in section 7.1, the height at which the cosmic ray flux interacts  
 2677 in the atmosphere is not known on an event-by-event basis. The production  
 2678 height can vary from the Earth’s surface to  $\sim 50\text{km}$  above that. The SK-only  
 2679 analysis methodology (described in section 7.2) for including the uncertainty  
 2680 on the production height is to include variations from the Honda model when  
 2681 pre-calculating the oscillation probabilities prior to the fit. This technique is not  
 2682 possible for this analysis which uses continuous oscillation parameters that can  
 2683 not be known prior to the fit. Consequently, an analytical averaging technique  
 2684 was developed in [208]. The author of this thesis was not responsible for the  
 2685 derivation of the technique but has performed the implementation and validation  
 2686 of the technique for this analysis alone.

2687 Using the 20 production heights per Monte Carlo neutrino event, provided  
 2688 as 5% percentiles from the Honda flux model, a production height distribution  
 2689  $p_j(h|E_\nu, \cos \theta_Z)$  is built for each neutrino flavour  $j = \nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$ . In practice, a  
 2690 histogram is filled with 20 evenly spaced bins in production height  $h$  between  
 2691 0 and 50km. The neutrino energy and cosine zenith binning of the histogram  
 2692 are the same as that provided in section 7.2. The average production height,  
 2693  $\bar{h} = \int dh \frac{1}{4} \sum_j p_j(h|E_\nu, \cos(\theta_Z))$ , is calculated. This assumes a linear average over  
 2694 the four flavours of neutrino which are considered to be generated in cosmic ray  
 2695 showers. The production height binning of this histogram is then translated into  
 2696  $\delta t(h) = t(\bar{h}) - t(h)$ , where  $t(h)$  is the distance travelled along the trajectory.

2697 For the  $i^{\text{th}}$  traversed layer, the transition amplitude,  $D_i(t_{i+1}, t_i)$ , is computed.  
 2698 The time-ordered product of these is then used as the overall transition amplitude  
 2699 via

$$A(t_{n+1}, t_0) = D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1) D_0(t_1, t_0), \quad (7.7)$$

2700 where,

$$\begin{aligned} D_n(t_{n+1}, t_n) &= \exp[-iH_n(t_{n+1} - t_n)] \\ &= \sum_{k=1}^3 C_k \exp[ia_k(t_{n+1} - t_n)] \end{aligned} \quad (7.8)$$

2701 is expressed as a diagonalised time-dependent solution to the Schrodinger  
 2702 equation. The 0<sup>th</sup> layer is the propagation through the atmosphere and is the  
 2703 only term that depends on the production height. Using the substitution  $t_0 =$   
 2704  $t(\bar{h}) - \delta t(h)$ , it can be shown that

$$D_0(t_1, t_0) = D_0(t_1, \bar{h})D_0(\delta t). \quad (7.9)$$

2705 Thus Equation 7.7 becomes

$$\begin{aligned} A(t_{n+1}, t_0) &= D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1)D_0(t_1, \bar{h})D(\delta t) \\ &= A(t_{n+1}, \bar{h}) \sum_{k=1}^3 C_k \exp[ia_k \delta t], \\ &= \sum_{k=1}^3 B_k \exp[ia_k \delta t]. \end{aligned} \quad (7.10)$$

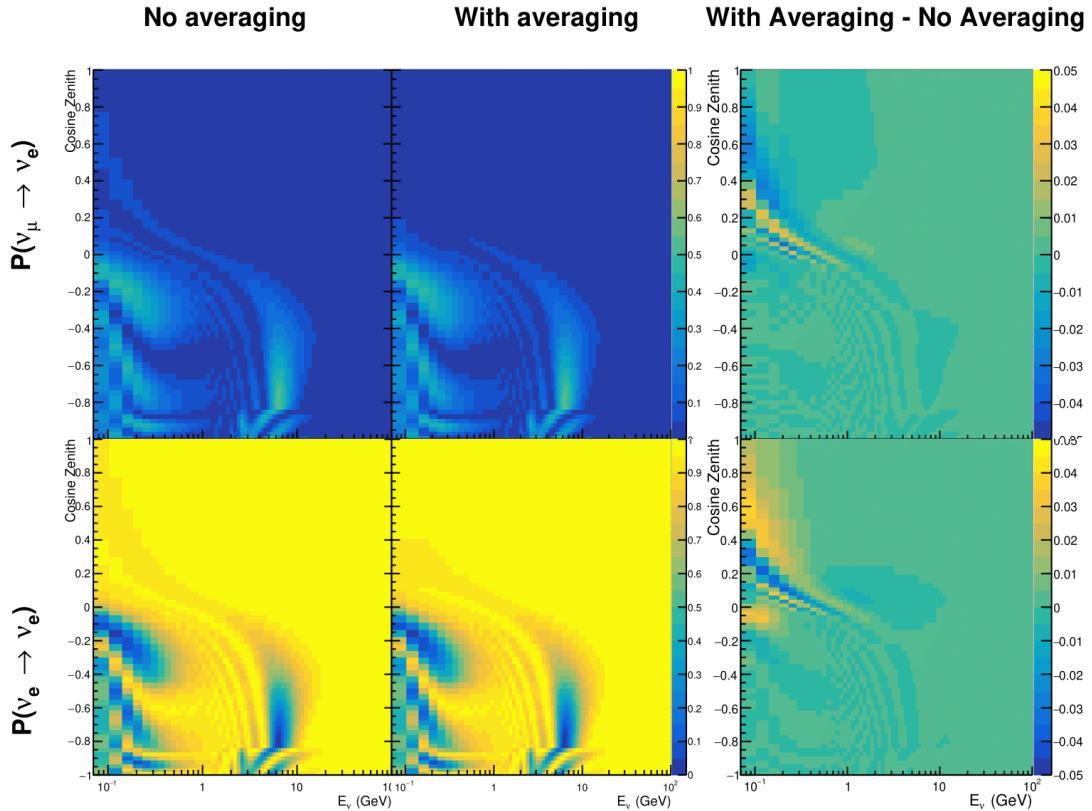
2706 The oscillation probability averaged over production height is then calculated  
 2707 as

$$\begin{aligned} \bar{P}(\nu_j \rightarrow \nu_i) &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) P(\nu_j \rightarrow \nu_i) \\ &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) A(t_{n+1}, t_0) A^*(t_{n+1}, t_0) \\ &= \sum_{km} (B_k)_{ij} (B_m)^*_{ij} \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) \exp[i(a_k - a_m) \delta t] \end{aligned} \quad (7.11)$$

2708 It is important to note that the exact value of  $\bar{h}$  used does not matter as the  
 2709 values of  $\delta t$  would change to compensate for any modification to the value of  $\bar{h}$ .

2710 In practice, implementation in CUDAProb3 [211] is relatively straightforward  
 2711 as the majority of these terms are already calculated in the standard oscillation  
 2712 calculation. Figure 7.19 illustrates the results of the production height averaging.

2713 As expected, the main effect is observed in the low-energy downward-going  
 2714 and horizontal-going events. Upward-going events have to travel the radius  
 2715 of the Earth,  $R_E = 6371\text{km}$ , where the production height uncertainty is a small  
 2716 fraction of the total path length.



**Figure 7.19:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where a fixed production height of 25km is used. The middle column gives the probabilities where the production height is analytically averaged. The right column illustrates the difference in oscillation probability between the two different techniques.

# 8

2717

2718

## Oscillation Analysis

2719 Using the samples and systematics defined in chapter 6, this chapter documents  
2720 a simultaneous beam and atmospheric oscillation analysis from the T2K and SK  
2721 experiments. The MaCh3 Bayesian MCMC framework introduced in chapter 4  
2722 is used for all studies performed within this thesis.

2723 The MaCh3 framework used throughout this thesis has been validated through  
2724 many tests. The code which handles the beam far detector samples was developed  
2725 by the author and validated by comparison to the 2020 T2K analysis [74]. The  
2726 sample event rates and likelihood evaluations of beam samples generated by  
2727 the framework used within this thesis were compared to those from the T2K  
2728 analysis by the author of this thesis. Variations of the sample predictions were  
2729 compared at  $\pm 1\sigma$  and  $\pm 3\sigma$  and good agreement was found in all cases. A similar  
2730 study, led by Dr. C. Wret was used to validate the near detector portion of  
2731 the code [201]. The implementation of the atmospheric samples within MaCh3  
2732 was completed and cross-checked by the author of this thesis against the P-  
2733 Theta framework (Introduced in section 3.2). Both fitters are provided with  
2734 the same inputs and can therefore cross-validate each other. These validations  
2735 compared the event rate and likelihood calculation. Documentation of all the  
2736 above validations can be found in [201].

## 2737 8.1 Monte Carlo Prediction

2738 Using the three sets of dial values (generated, pre-BANFF, and post-BANFF tunes)  
 2739 defined in subsection 6.4.3, the predicted event rates for each sample are defined  
 2740 in Table 8.1. The oscillated (AsimovA defined in Table 2.2) and un-oscillated  
 2741 event rates are calculated for each tune.

Sample	Total Predicted Events					
	Generated		Pre-fit		Post-fit	
	Osc	UnOsc	Osc	UnOsc	Osc	UnOsc
SubGeV- <i>elike</i> -0dcy	7121.0	7102.6	6556.8	6540.0	7035.2	7015.7
SubGeV- <i>elike</i> -1dcy	704.8	725.5	693.8	712.8	565.7	586.0
SubGeV- <i>mulike</i> -0dcy	1176.5	1737.2	1078.6	1588.1	1182.7	1757.1
SubGeV- <i>mulike</i> -1dcy	5850.7	8978.1	5351.7	8205.1	5867.0	9009.9
SubGeV- <i>mulike</i> -2dcy	446.9	655.2	441.6	647.7	345.9	505.6
SubGeV- <i>pi0like</i>	1438.8	1445.4	1454.9	1461.1	1131.1	1136.2
MultiGeV- <i>elike</i> -nue	201.4	195.6	201.1	195.3	202.6	196.7
MultiGeV- <i>elike</i> -nuebar	1141.5	1118.3	1060.7	1039.5	1118.5	1095.7
MultiGeV- <i>mulike</i>	1036.7	1435.8	963.1	1334.1	1015.2	1405.9
MultiRing- <i>elike</i> -nue	1025.1	982.2	1026.8	984.3	1029.8	986.4
MultiRing- <i>elike</i> -nuebar	1014.8	984.5	991.0	962.0	1008.9	978.5
MultiRing- <i>mulike</i>	2510.0	3474.4	2475.6	3425.8	2514.6	3480.4
MultiRingOther-1	1204.5	1279.1	1205.8	1280.3	1207.4	1281.0
PCStop	349.2	459.2	338.4	444.7	346.8	456.1
PCThru	1692.8	2192.5	1661.5	2149.8	1689.2	2187.8
UpStop-mu	751.2	1295.0	739.7	1271.6	750.4	1293.0
UpThruNonShower-mu	2584.4	3031.6	2577.9	3019.4	2586.8	3034.0
UpThruShower-mu	473.0	488.6	473.2	488.7	473.8	489.4
FHC1Rmu	328.0	1409.2	301.1	1274.7	345.1	1568.0
RHC1Rmu	133.0	432.3	122.7	396.2	135.0	443.9
FHC1Re	84.6	19.2	77.4	18.2	93.7	19.7
RHC1Re	15.7	6.4	14.6	6.1	15.9	6.3
FHC1Re1de	10.5	3.2	10.3	3.1	8.8	2.9

**Table 8.1:** The Monte Carlo predicted event rate of each far detector sample used within this analysis. Three model parameter tunes are considered, as defined in subsection 6.4.3. Un-oscillated and oscillated predictions are given, where the oscillated predictions assume Asimov A oscillation parameters provided in Table 2.2.

2742 Generally, the samples which target CCQE interaction modes observe a  
 2743 decrease in prediction when using the pre-fit dial values. This is in accordance  
 2744 with the Monte Carlo being produced assumed  $M_A^{QE} = 1.21\text{GeV}$  [161] whilst

2745 the pre-fit dial value is set to  $M_A^{QE} = 1.03\text{GeV}$  as suggested by [199]. Further-  
 2746 more, the predicted event rates of samples that target CCRES interaction modes  
 2747 are significantly reduced when considering the post-BANFF fit. This follows  
 2748 the observations in subsection 6.4.3. The strength of the accelerator neutrino  
 2749 experiment can be seen in the remarkable difference between the oscillated and  
 2750 unoscillated predictions in the FHC1Rmu and RHC1Rmu samples. There is a very  
 2751 clear decrease in the expected event rate between the oscillated and un-oscillated  
 2752 predictions which is not as obvious as in the atmospheric samples. This is due to  
 2753 the fact that the beam energy is tuned to the maximum disappearance probability  
 2754 which is not the case for the naturally generated atmospheric neutrinos.

## 2755 8.2 Likelihood Calculation

### 2756 DB: Should this section be moved?

2757 This analysis performs a joint oscillation parameter fit of the ND280 beam  
 2758 samples, the T2K far detector beam samples, and the SK atmospheric samples.

2759 Once the Monte Carlo predictions of each beam and atmospheric sample  
 2760 have been built (see chapter 6) a likelihood needs to be constructed. This is done  
 2761 by comparing the binned Monte Carlo prediction to binned data. The Monte  
 2762 Carlo prediction is calculated at a particular point,  $\vec{\theta}$ , in the model parameter  
 2763 space such that  $N_i^{MC} = N_i^{MC}(\vec{\theta})$ , where  $N_i$  represents the bin content of the  
 2764  $i^{th}$  bin. The data and Monte Carlo spectra are represented by  $N_i^D$  and  $N_i^{MC}$ ,  
 2765 respectively. The bin contents for the beam near detector, beam far detector and  
 2766 atmospheric samples are denoted with  $ND$ ,  $FD$ , and  $Atm$ , respectively. The  
 2767 binning index,  $i$ , runs over all the bins within a sample. Taking the FHC1Rmu far  
 2768 detector sample as an example, the binning index runs over all the reconstructed  
 2769 neutrino energy bins. The likelihood calculation between the data and the Monte  
 2770 Carlo prediction for a particular bin follows a Poisson distribution, where the  
 2771 data is treated as a fluctuation of the simulation.

2772 The data can consist of either real data or an ‘Asimov’ Monte Carlo prediction,  
 2773 which is typically used for sensitivity studies and denoted ‘Asimov data’. The

process for building Asimov data is as follows. The Monte Carlo prediction is reweighted using a particular set of oscillation parameters (potentially those listed in Table 2.2) and systematic parameter tune. The resulting spectra for each sample is then defined to be the Asimov data for that sample. Whilst this results in unphysical non-integer data predictions, it eliminates statistical fluctuations from the data. Therefore, the results of a fit to Asimov data should not include any biases from statistical fluctuations. Furthermore, these results should produce posterior probability distributions consistent with the parameters which were used to make the data prediction. That is to say, the fit results should return the known parameters. Any biases seen would be attributed to correlations between each oscillation parameter and correlations between oscillation and systematic parameters. Consequently, Asimov fit results present the maximum precision at which the oscillation parameters could be measured to.

Following the T2K analysis presented in [74], the likelihood contribution for the near detector samples also includes a Monte Carlo statistical uncertainty term, derived from the Barlow and Beeston statistical treatment [217, 218]. In addition to treating the data as a Poisson fluctuation of the Monte Carlo prediction, it includes a contribution to the likelihood that which treats the generated Monte Carlo prediction as a statistical fluctuation of the actual true simulation assuming an infinite amount of statistics had been created. The technical implementation of this additional likelihood term is documented in [190] and briefly summarised as follows. The term is defined as,

$$\frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2}, \quad (8.1)$$

where  $\beta_i$  represents a scaling parameter for the  $i^{th}$  bin that relates the bin content for the amount of Monte Carlo actually generated  $N_i^{MC}$  to the bin content if an infinite amount of Monte Carlo statistics had been generated  $N_{i,true}^{MC}$ , such that  $N_{i,true}^{MC} = \beta_i \times N_i^{MC}$ . In the case where a sufficient amount of Monte Carlo statistics had been generated,  $\beta_i = 1$ . An analytical solution for  $\beta_i$  is given in

[190]. Additionally,  $\sigma_{\beta_i} = \sqrt{\sum_i w_i^2} / N_i^{MC}$  where  $\sqrt{\sum_i w_i^2}$  represents the sum of the square of the weights of the Monte Carlo events which fall into bin  $i$ .

An additional contribution to the likelihood comes from the variation of the systematic model parameters. For those parameters with well-motivated uncertainty estimates, a covariance matrix,  $V$ , describes the prior knowledge of each parameter as well as any correlations between the parameters. Due to a technical implementation, a single covariance matrix describes each “block” of model parameters, e.g. beam flux systematics. The covariance matrix associated with the  $k^{th}$  block is denoted  $V^k$ . This substitution results in  $\vec{\theta} = \sum_k^{N_b} \vec{\theta}^k$  and  $V = \sum_k^{N_b} V^k$  where  $N_b$  denotes the number of blocks. A single covariance matrix is provided for: the oscillation parameters, the beam flux parameters, the atmospheric flux parameters, the neutrino interaction systematics, the near detector parameters, the beam far detector systematics, and the atmospheric far detector systematics. All of these are detailed in section 6.4. The number of parameters in the  $k^{th}$  block is defined as  $n(k)$ .

The equation for the likelihood  $\mathcal{L}$  includes all the terms discussed above.

It is defined as,

$$\begin{aligned}
 -\ln(\mathcal{L}) = & \\
 & \sum_i^{\text{NDbins}} N_i^{\text{ND},MC}(\vec{\theta}) - N_i^{\text{ND},D} + N_i^{\text{ND},D} \times \ln \left[ N_i^{\text{ND},D} / N_i^{\text{ND},MC}(\vec{\theta}) \right] + \frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2} \\
 & + \sum_i^{\text{FDbins}} N_i^{\text{FD},MC}(\vec{\theta}) - N_i^{\text{FD},D} + N_i^{\text{FD},D} \times \ln \left[ N_i^{\text{FD},D} / N_i^{\text{FD},MC}(\vec{\theta}) \right] \\
 & + \sum_i^{\text{Atmbins}} N_i^{\text{Atm},MC}(\vec{\theta}) - N_i^{\text{Atm},D} + N_i^{\text{Atm},D} \times \ln \left[ N_i^{\text{Atm},D} / N_i^{\text{Atm},MC}(\vec{\theta}) \right] \\
 & + \frac{1}{2} \sum_k^{N_b} \sum_i^{n(k)} \sum_j^{n(k)} (\vec{\theta}^k)_i (V^k)_{ij}^{-1} (\vec{\theta}^k)_j.
 \end{aligned} \tag{8.2}$$

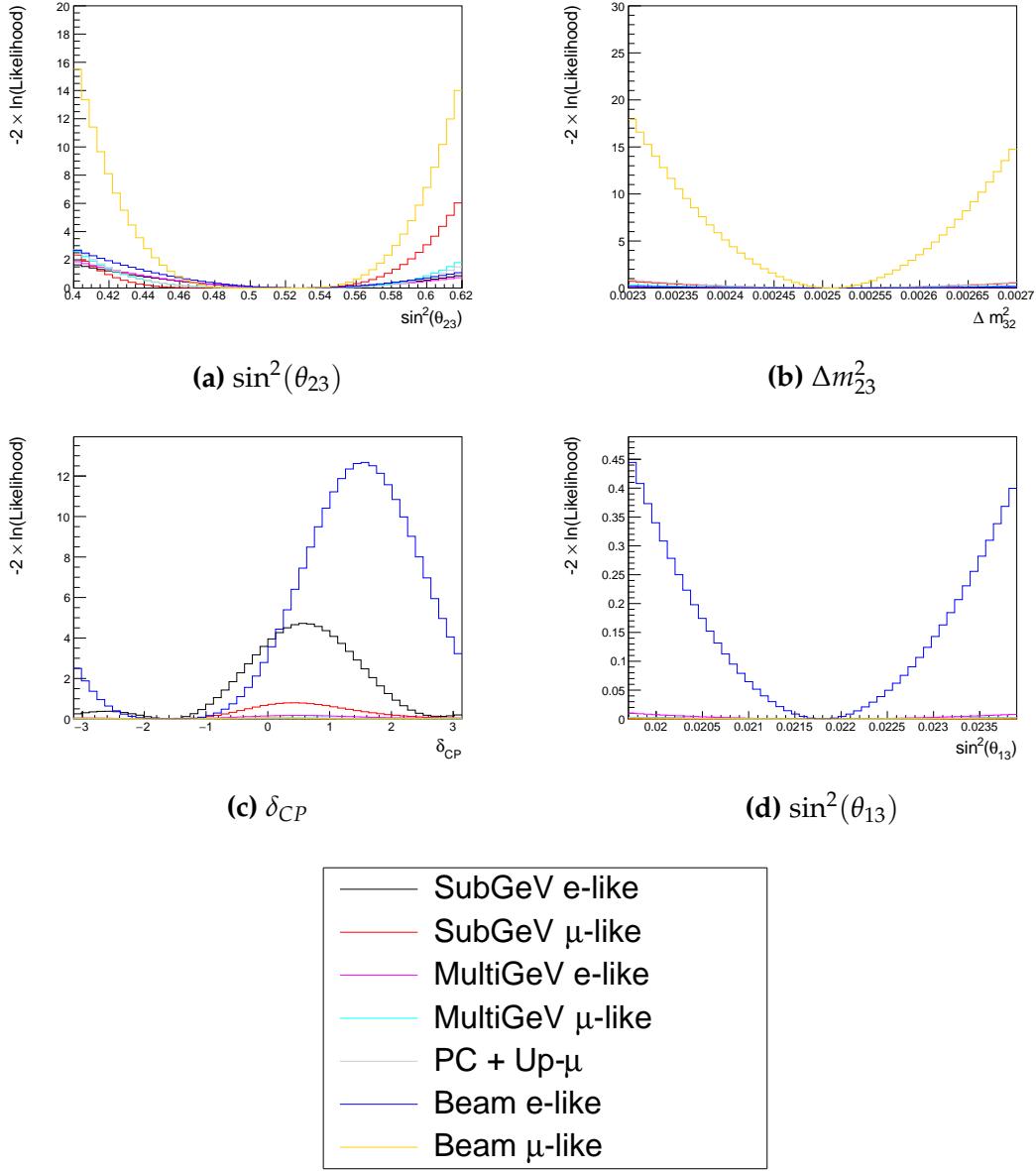
The negative log-likelihood value is determined at each step of the MCMC to build the posterior distribution defined in chapter 4. This value is minimised when the Monte Carlo prediction tends towards the data spectrum.

### 8.2.1 Likelihood Scans

Using the definition of the likelihood presented in section 8.2, the response of each sample to a variation of a particular parameter can be studied. Figure 8.1 presents the variation of all the samples (beam and atmospheric) at the far detector to the oscillation parameters of interest:  $\delta_{CP}$ ,  $\sin^2(\theta_{13})$ ,  $\sin^2(\theta_{23})$ , and  $\Delta m_{32}^2$ . These plots are colloquially called ‘likelihood scan’ (or ‘log-likelihood scans’). The process of making these plots is as follows. An Asimov data set (following technique detailed in section 8.2) is built using the AsimovA oscillation parameters and pre-fit systematic tune. The Monte Carlo is then reweighted using the value of the oscillation parameter at each point on the x-axis of the scan. The likelihood is then calculated between the Asimov data and Monte Carlo prediction and plotted. This process identifies which samples drive the determination of the oscillation parameters in the joint fit.

Due to the caveat of fixed systematic parameters and the correlations between oscillation parameters being ignored when creating these likelihood scans, the value of  $\chi^2 = 1$  (or  $-2 \times \ln(\text{Likelihood}) = 1$ ) does not equate to the typical  $1\sigma$  sensitivity. However, it does give an indication of which samples respond the strongest to a variation in a particular oscillation parameter. The point at which the likelihood tends to zero illustrates the value of the parameter used to build the Asimov data prediction.

The sensitivity to  $\sin^2(\theta_{23})$  is mostly dominated by the beam muon-like samples. The response of an individual atmospheric sample is small but non-negligible such that the summed response over all atmospheric samples becomes comparable to that of the muon-like beam samples. Consequently, the sensitivity of the joint fit to  $\sin^2(\theta_{23})$  would be expected to be greater than the beam-only analysis. The only sample which respond to the  $\sin^2(\theta_{13})$  oscillation parameter is the electron-like beam samples. Consequently, no increase in sensitivity beyond that of the T2K-only analysis would be expected from the joint fit. Regardless, the sensitivity of the beam sample is significantly weaker than the external reactor constraint so prior knowledge will dominate any measurement that is included



**Figure 8.1:** The response of the likelihood, as defined in section 8.2, illustrating the response of the samples to a variation of an oscillation parameter.

within this thesis. The  $\Delta m_{21}^2$  and  $\sin^2(\theta_{12})$  parameters are not considered as there is simply no sensitivity in any sample considered within this analysis. The response to  $\Delta m_{32}^2$  is completely dominated by the beam muon-like samples. This is because the beam neutrino energy can be specifically tuned to match the maximal disappearance probability. Despite this, improvements to the  $|\Delta m_{32}^2|$  sensitivity may be expected due to additional mass hierarchy determination

2857 added by the atmospheric samples.

2858 Two-dimensional scans of the appearance ( $\sin^2(\theta_{13})-\delta_{CP}$ ) and disappearance  
2859 ( $\sin^2(\theta_{23})-\Delta m_{32}^2$ ) parameters are illustrated in Figure 8.2 and Figure 8.3, respec-  
2860 tively. The caveat of fixed systematic parameters and correlations between other  
2861 oscillation parameters being neglected still apply.

2862 The appearance log-likelihood scans show the distinct difference in how  
2863 the beam and atmospheric samples respond. The beam samples have an ap-  
2864 proximately constant width of the  $2\sigma$  and  $3\sigma$  contours, throughout all ranges  
2865 of  $\delta_{CP}$ . The response of the atmospheric samples to  $\sin^2(\theta_{13})$  is very strongly  
2866 correlated to the value of  $\delta_{CP}$  being evaluated, with the strongest constraints  
2867 around  $\delta_{CP} \sim 1$ . Consequently, this difference allows some of the degeneracy in  
2868 a beam-only fit to be broken. Comparing the beam-only and joint fit likelihood  
2869 scans, the  $2\sigma$  continuous contour in  $\delta_{CP}$  for beam samples is broken when the  
2870 atmospheric samples are added. This may result in a stronger sensitivity to  
2871  $\delta_{CP}$ . Similarly, the width of the  $3\sigma$  contours also becomes dependent upon the  
2872 value of  $\delta_{CP}$ . Whilst these are encouraging results for the joint fit, these are  
2873 not sensitivity measurements as the systematic parameters are fixed and the  
2874 correlation between oscillation parameters is neglected. An interesting point to  
2875 note is that the atmospheric samples have little sensitivity to  $\sin^2(\theta_{13})$  on their  
2876 own, as evidenced in Figure 8.1, but can improve sensitivity to the parameter  
2877 when combined within the simultaneous fit.

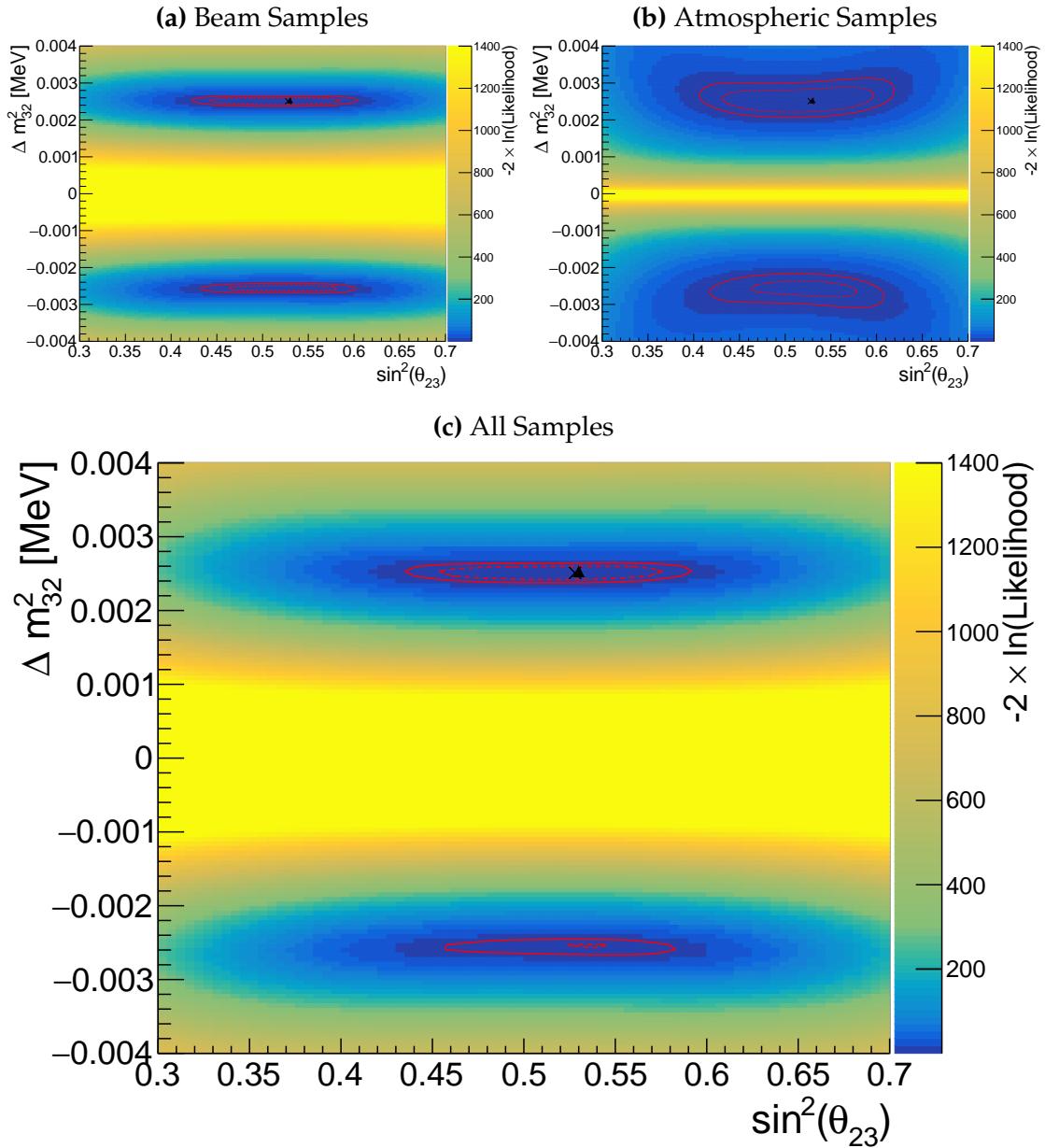
2878 The response of the atmospheric samples in Figure 8.2 shows an interesting  
2879 behaviour when considering the application of the reactor constraint. At higher  
2880 values of  $\sin^2(\theta_{13})$ , two lobes appear around  $\delta_{CP} \sim -\pi/2$  and  $\delta_{CP} \sim 2.4$ . If this  
2881 distribution was projected onto the  $\delta_{CP}$  axis, these lobes would mean the posterior  
2882 distribution would have a significant dip between these values. However, the  
2883 region of  $\sin^2(\theta_{13})$  near the reactor constraint ( $\sin^2(\theta_{13}) = (2.18 \pm 0.08) \times 10^{-2}$ )  
2884 is flatter across the range of  $\delta_{CP}$ . Therefore, if we were to project only this  
2885 region onto the  $\delta_{CP}$  axis, the dip between the peaks would not be as significant.  
2886 If this behaviour was to be seen in the results of a fit, these marginalisation

2887 effects would actually conspire to reduce the sensitivity to  $\delta_{CP}$  if the reactor  
2888 constraint was to be applied.

2889 The disappearance log-likelihood scans in  $\sin^2(\theta_{23})$ - $\Delta m_{32}^2$  space (Figure 8.3)  
2890 show the expected behaviour when considering the one-dimensional scans  
2891 already discussed. The uncertainty on the width of  $|\Delta m_{32}^2|$  is mostly driven  
2892 by the beam-only sensitivities. However, the width of this contour in the  
2893 inverted mass region ( $\Delta m_{32}^2 < 0$ ) is significantly reduced due to the ability  
2894 of the atmospheric samples to select the correct (normal) mass hierarchy. The  
2895 width of the uncertainty in  $\sin^2(\theta_{23})$  is also reduced compared to the beam-only  
2896 sensitivities, with a further decrease in the inverted hierarchy region due to the  
2897 better mass hierarchy determination.



**Figure 8.2:** Two-dimensional log-likelihood scan of the appearance ( $\sin^2(\theta_{13})$ - $\delta_{CP}$ ) parameters showing the response of the beam samples (top left), atmospheric samples (top right) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are known to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.



**Figure 8.3:** Two-dimensional log-likelihood scan of the disappearance ( $\sin^2(\theta_{23})$ - $\Delta m_{32}^2$ ) parameters showing the response of the beam samples (top left), atmospheric samples (top right) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are known to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.

The likelihood scans illustrated thus far only consider the sensitivity of this analysis for a fixed set of true oscillation parameters, namely Asimov A defined in Table 2.2. Whilst computational infeasible to run many fits at different parameter sets, it is possible to calculate the likelihood response to different Asimov data sets. Figure 8.4 and Figure 8.5 illustrate how the sensitivity changes for differing true values of  $\delta_{CP}$  and  $\sin^2(\theta_{23})$ , respectively. For both of these plots, the other oscillation parameters are fixed at their Asimov A values. Consequently, the caveat of fixed systematic parameters and correlations between other oscillation parameters being neglected still applies.

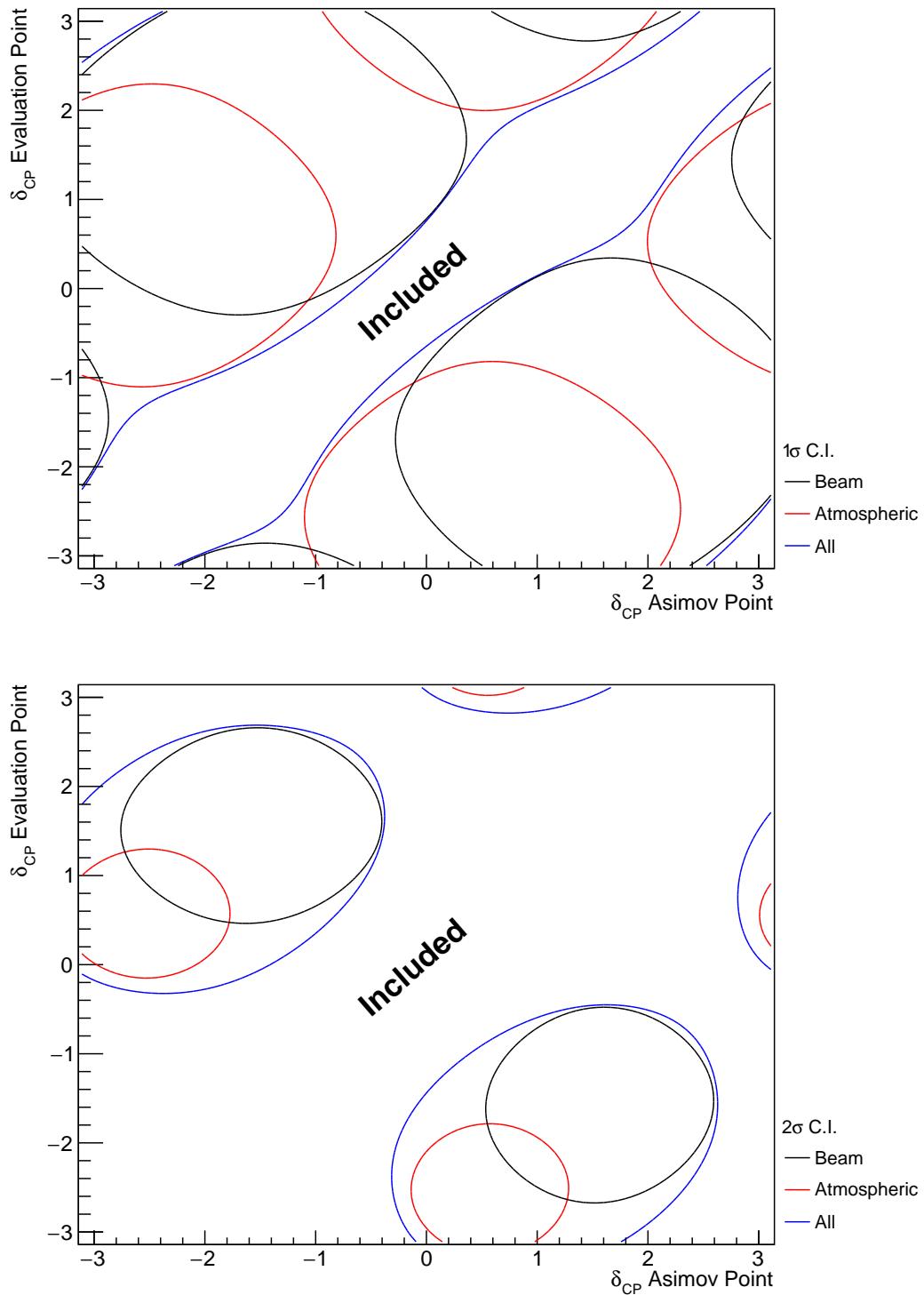
To explain how these plots are made, consider Figure 8.4. This plot is built by considering multiple one-dimensional log-likelihood scans, each creating an Asimov data with the value of  $\delta_{CP}$  taken from the x-axis. The likelihood to a particular Asimov data set is calculated after reweighting the Monte Carlo prediction to each value of  $\delta_{CP}$  on the y-axis.

Figure 8.4 illustrates the sensitivity to  $\delta_{CP}$ . Notably, the  $1\sigma$  intervals contain regions in the off-diagonal for which the beam and atmospheric samples have broken and discontinuous contours. This indicates that there are regions of  $\delta_{CP}$  which are degenerate. For example, for the x-axis value of  $\delta_{CP} = 0$ , the beam samples sensitivity would include two discontinuous regions excluded from the  $1\sigma$  interval:  $\delta_{CP} \sim 0$  and  $\delta_{CP} \sim \pi$ . The offset in  $\delta_{CP}$  between the beam and atmospheric samples allows the joint fit to have increased sensitivity in these regions. Consequently, the difference between the beam-only and joint beam-atmospheric fit should be studied using multiple Asimov data sets.

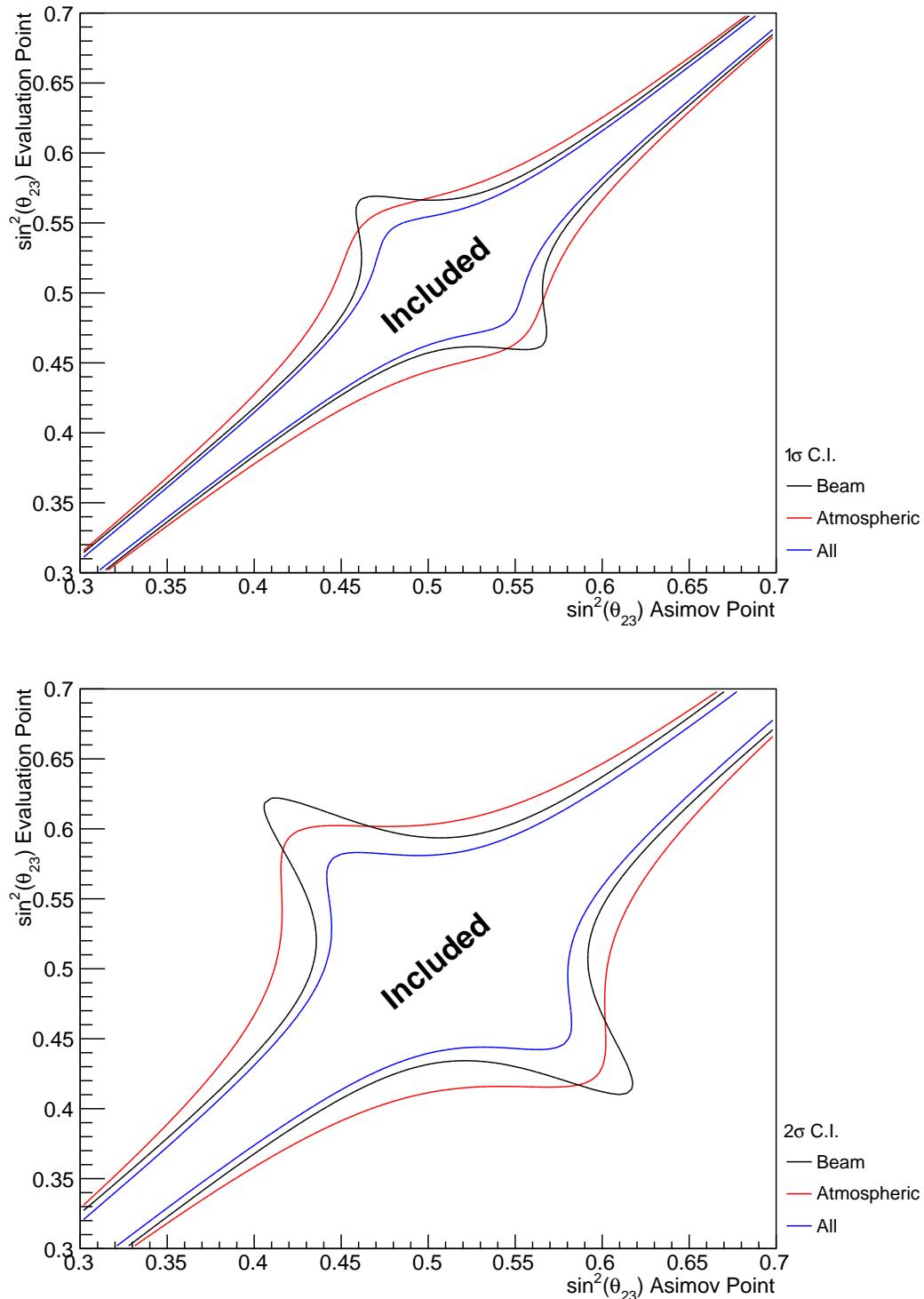
Despite the increased sensitivity at  $1\sigma$ , the  $2\sigma$  intervals from the joint fit are more similar to the two independent sensitivities and the off-diagonal degeneracies mostly remain. This indicates that the joint fit has the strength to aid parameter determination but can not entirely break the degeneracies in  $\delta_{CP}$  at higher confidence levels.

Figure 8.5 illustrates a similar analysis as above, although the value of  $\sin^2(\theta_{23})$  is varied and  $\delta_{CP}$  is fixed to the Asimov A parameter value. Due to the beam

parameters and baseline being tuned to specifically target this oscillation parameter, the average sensitivity of the beam samples is stronger than the atmospheric samples. However, the degeneracy around maximal mixing ( $\sin^2(\theta_{23}) = 0.5$ ) is significantly more peaked in the beam samples compared to the atmospheric samples. This behaviour is strengthened when considering the  $2\sigma$  intervals, to the point where two distinct discontinuous regions of the  $2\sigma$  intervals exist around the Asimov point  $\sin^2(\theta_{23}) \sim 0.41, 0.6$ . Given the caveat of only considering likelihood scans, the joint analysis would mostly eliminate the discontinuous intervals in these regions. This means that the joint fit could feasibly have an increased preference for the correct octant hypothesis.

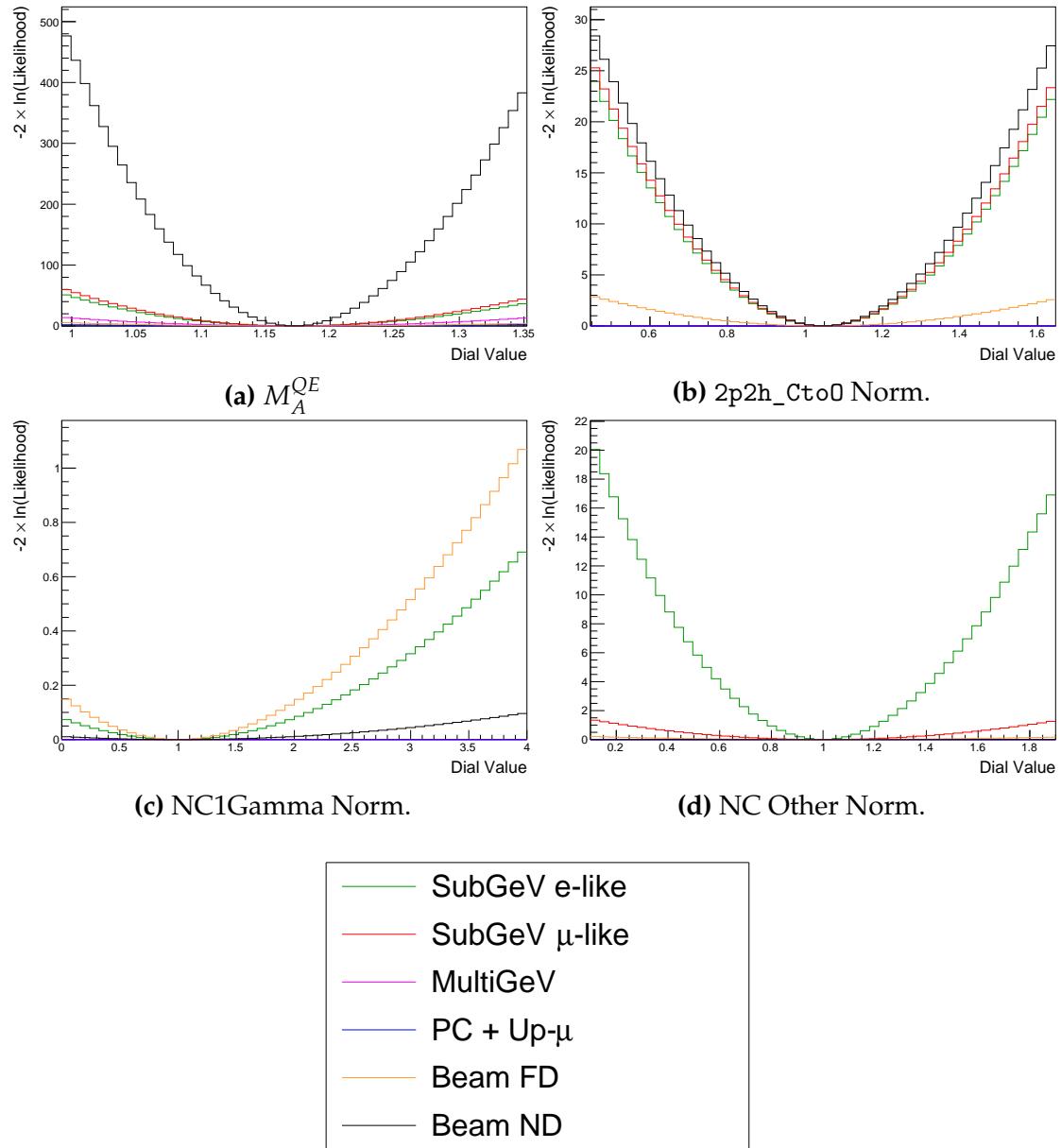


**Figure 8.4:** A series of one-dimensional likelihood scans over  $\delta_{CP}$ , where an Asimov data set is built for each value of  $\delta_{CP}$  on the x-axis and the likelihood is evaluated for each value of  $\delta_{CP}$  on the y-axis. The diagonal represents the minimum log-likelihood and defines the region included within the  $1\sigma$  (Top) and  $2\sigma$  (Bottom) confidence intervals. The beam (black) and atmospheric (red) samples are individually plotted and the joint fit (blue) is the sum of the two.



**Figure 8.5:** A series of one-dimensional likelihood scans over  $\sin^2(\theta_{23})$ , where an Asimov data set is built for each value of  $\sin^2(\theta_{23})$  on the x-axis and the likelihood is evaluated for each value of  $\sin^2(\theta_{23})$  on the y-axis. The diagonal represents the minimum log-likelihood and defines the region included within the  $1\sigma$  (Top) and  $2\sigma$  (Bottom) confidence intervals. The beam (black) and atmospheric (red) samples are individually plotted and the joint fit (blue) is the sum of the two.

Alongside oscillation parameters (Figure 8.1), the sensitivity to systematic parameters can also be studied for the joint fit. As some of these parameters are correlated between the beam and atmospheric events, the response of the atmospheric samples can modify the constraint. This means the systematics can have additional constraints than what they would from a beam-only analysis. Therefore, the response from the beam and the atmospheric samples to various systematic parameters has been compared in Figure 8.6. The Asimov data set has been created using the AsimovA oscillation parameter and the pre-fit systematic tune. For example, the systematic parameter controlling the effective axial mass coupling in CCQE interactions,  $M_A^{QE}$ , is clearly dominated by the ND constraint. An example where the response of the atmospheric sample is approximately similar to the near detector constraint is the 2p2h\_Cto0 normalisation systematic. This systematic models the scaling of the 2p2h interaction cross-section on a carbon target to an oxygen target. There are also systematics which have no near detector constraint. For example, the systematic parameters which describe the normalisation of the NC1Gamma and NCOther interaction modes. The atmospheric samples are significantly more sensitive to these systematics than the beam samples due to their similar interaction contributions but relatively higher statistics (Table 8.1). As an example of how the atmospheric samples can help constrain systematic parameters used within the T2K-only analysis, these NC background events in beam electron-like samples will be considerably more constrained with the additional sensitivity of atmospheric samples. This would be expected to reduce the overall uncertainty of the beam electron-like event rates in the joint analysis compared to the beam-only studies. This could modify the sensitivity of the beam samples due to the more constrained background events.



**Figure 8.6:** The response of the likelihood, as defined in section 8.2, illustrating the response of the samples to the various cross-section systematic parameters.

### 2963 8.3 Sensitivities

2964 The sensitivities of the joint T2K and SK oscillation analysis are presented in  
2965 the form of Asimov fits. This technique builds an Asimov data set (follow-  
2966 ing section 8.2) using the AsimovA oscillation parameters and post-BANFF  
2967 systematic tune.

2968 In practice, the Asimov fits presented within this analysis are modified from  
2969 the above definition. An Asimov prediction of both beam and atmospheric far  
2970 detector samples is fit whilst the true data is used for near detector samples.  
2971 The Asimov predictions at the far detector are built using the BANFF tuning (as  
2972 discussed in section 3.2). These modifications mean that the results are equivalent  
2973 to performing a far detector Asimov fit using inputs from the BANFF data fit.  
2974 Consequently, this allows the results to be cross-checked with the results from  
2975 the P-Theta analysis. The comparison has been performed and is documented in  
2976 [219]. No significant discrepancies were found between the fitters.

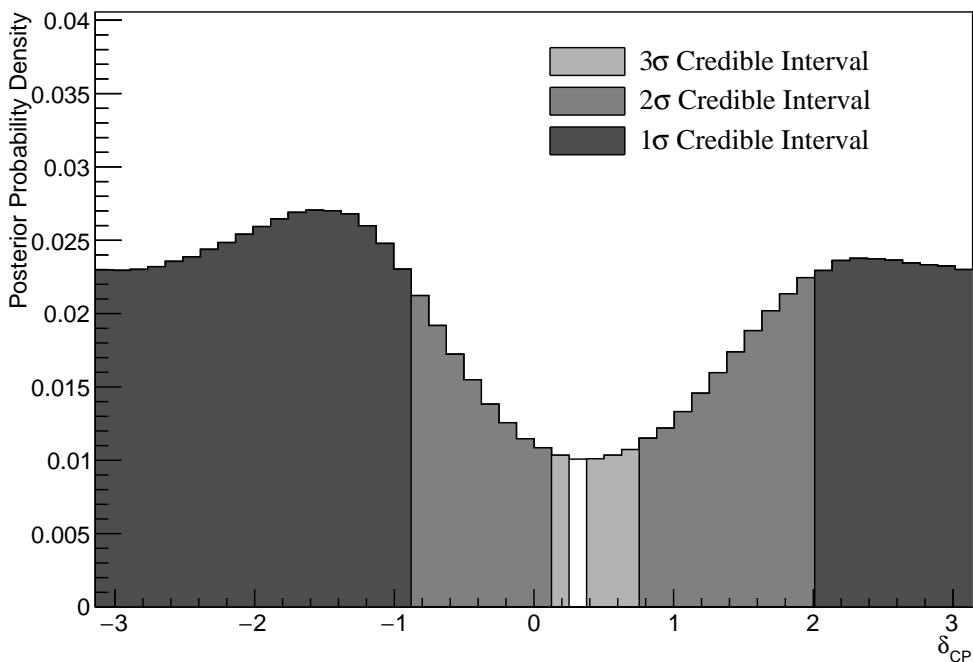
2977 This section proceeds with the following studies. Firstly, the sensitivity of  
2978 the atmospheric samples after the T2K cross-section has been applied to the low-  
2979 energy events is detailed in subsection 8.3.1. This includes studying the choice of  
2980 applying the 2020 PDG reactor constraint [75] to the atmospheric samples, which  
2981 is documented in subsection 8.3.2. Additionally, the effect of applying the near-  
2982 detector constraints onto the atmospheric samples is discussed in subsection 8.3.3.  
2983 The main result is the sensitivity of the simultaneous beam and atmospheric fit.  
2984 The sensitivities, both with and without the application of the reactor constraint,  
2985 are presented in subsection 8.3.4 and subsection 8.3.5, respectively. To indicate  
2986 the benefit of the joint analysis, the sensitivities are compared to the 2020 T2K  
2987 sensitivities [74, 182] in subsection 8.3.6 and subsection 8.3.7. As shown in  
2988 subsection 8.2.1, the response of the beam and atmospheric samples change  
2989 depending upon the true set of oscillation parameters assumed. Therefore,  
2990 subsection 8.3.8 documents the sensitivities at an alternative oscillation parameter

<sup>2991</sup> set. It is important to note that these results have been published at the Neutrino  
<sup>2992</sup> 2022 conference on behalf of the T2K and SK collaborations [108].

### 2993 8.3.1 Atmospheric-Only Sensitivity Without Reactor Constraint

2994 This section presents the results of an Asimov fit using samples from the near  
 2995 detector and only atmospheric samples from the far detector. The results are  
 2996 presented as one-dimensional or two-dimensional histograms which have been  
 2997 marginalised over all other parameters using the technique outlined in sub-  
 2998 section 4.3.1. Each histogram displays the posterior probability density and  
 2999 illustrates the credible intervals, calculated using the technique in subsection 4.3.2.  
 3000 For these fits in this subsection, a flat penalty term is used for  $\sin^2(\theta_{13})$  such the  
 3001 reactor constraint is not applied. The Asimov data is generated assuming the  
 3002 AsimovA oscillation parameter set defined in Table 2.2 and the post-BANFF  
 3003 systematic parameter tune.

Without Reactor Constraint, Both Hierarchies



**Figure 8.7:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$ , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3004 Figure 8.7 illustrates the posterior probability density for  $\delta_{CP}$ , marginalised  
 3005 over both hierarchies. If instead, only steps in the normal hierarchy were  
 3006 considered, the shape of the contours would change. The fit favours the known

oscillation parameter ( $\delta_{CP} = -1.601$ ) although the posterior probability is very flat through the range of  $-\pi < \delta_{CP} < -1$  and  $2 < \delta_{CP} < \pi$ . There is also a region around  $\delta_{CP} \sim 0.4$  which is disfavoured at  $2\sigma$ . This indicates that the SK samples can rule out some parts of the CP conserving parameter space reasonably well, near  $\delta_{CP} \sim 0.4$ , when the true value of  $\delta_{CP} \sim \pi/2$ .

### Without Reactor Constraint, Both Hierarchies



**Figure 8.8:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$ , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

The posterior probability density in  $\Delta m_{32}^2$  is given in Figure 8.8. This distribution includes steps in both the normal hierarchy (NH,  $\Delta m_{32}^2 > 0$ ) and the inverse hierarchy (IH,  $\Delta m_{32}^2 < 0$ ). The highest posterior probability density is found within the NH, which agrees with the known oscillation parameter value. However, all of the credible intervals span both of the hierarchies hypotheses. If instead, only steps in the normal hierarchy were considered, the shape of the contours would change. The known oscillation parameter is  $2.509 \times 10^{-3} \text{ eV}^2$ , which is contained within the  $1\sigma$  credible interval.

	LO ( $\sin^2 \theta_{23} < 0.5$ )	UO ( $\sin^2 \theta_{23} > 0.5$ )	Sum
NH ( $\Delta m_{32}^2 > 0$ )	0.17	0.40	0.58
IH ( $\Delta m_{32}^2 < 0$ )	0.13	0.29	0.42
Sum	0.31	0.69	1.00

**Table 8.2:** The distribution of steps in an SK atmospheric-only fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as  $B(\text{NH}/\text{IH}) = 1.37$  and  $B(\text{UO}/\text{LO}) = 2.24$ .

Following the discussion in subsection 4.3.3, the Bayes factor for hierarchy preference can be calculated by determining the fraction of steps that fall into the NH and the IH regions, as an equal prior is placed on both hypotheses. A similar calculation can be performed by calculating the fraction of steps which fall in the lower octant (LO,  $\sin^2 \theta_{23} < 0.5$ ) or upper octant (UO,  $\sin^2 \theta_{23} > 0.5$ ). The fraction of steps, broken down by hierarchy and octant, are given in Table 8.2. The Bayes factor for preferred hierarchy model is  $B(\text{NH}/\text{IH}) = 1.37$ . Table 4.1 states this value of the Bayes factor indicates a weak preference for the normal hierarchy model. The Bayes factor for choice of octant is  $B(\text{UO}/\text{LO}) = 2.24$ . This is also classified as a weak preference for the UO. Both of these show that the fit is returning the correct choice of models (NH and UO) for the known Asimov A oscillation parameters defined in Table 2.2.

The  $1\sigma$  credible intervals, broken down by hierarchy, and position in parameter space of the highest posterior probability density is given in Table 8.3. These are taken from the one-dimensional projections of the oscillation parameters, marginalised over all other parameters within the fit. For the known Asimov value of  $\delta_{CP} = -1.601$ , the  $1\sigma$  credible interval rules out a region between  $\delta_{CP} = -0.86$  and  $\delta_{CP} = 1.96$ , when marginalising over both hierarchies. The position of the highest posterior density is  $\delta_{CP} = -1.57$  which is clearly compatible with the known oscillation parameter value.

The sensitivity of the atmospheric samples to  $\sin^2(\theta_{13})$  is presented in Figure 8.9. The likelihood scans presented in Figure 8.1 suggest that the sensitivity

Parameter	Interval	HPD
$\delta_{CP}$ , (BH)	$[-\pi, -0.86], [1.96, \pi]$	-1.57
$\delta_{CP}$ , (NH)	$[-\pi, -0.86], [1.88, \pi]$	-1.57
$\delta_{CP}$ , (IH)	$[-\pi, -0.94], [1.96, \pi]$	-1.57
$\Delta m_{32}^2$ (BH) [ $\times 10^{-3}\text{eV}^2$ ]	$[-3.00, -2.50], [2.35, 3.15]$	2.65
$\Delta m_{32}^2$ (NH) [ $\times 10^{-3}\text{eV}^2$ ]	$[2.39, 3.04]$	2.64
$\Delta m_{32}^2$ (IH) [ $\times 10^{-3}\text{eV}^2$ ]	$[-3.15, -2.45]$	-2.70
$\sin^2(\theta_{23})$ (BH)	$[0.476, 0.59]$	0.542
$\sin^2(\theta_{23})$ (NH)	$[0.476, 0.59]$	0.554
$\sin^2(\theta_{23})$ (IH)	$[0.476, 0.59]$	0.542

**Table 8.3:** The position of the highest posterior probability density (HPD) and width of the  $1\sigma$  credible interval for the SK atmospheric-only fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

3042 to  $\sin^2(\theta_{13})$  will be small. This behaviour is also seen in the fit results, where the  
 3043 width of the  $1\sigma$  credible intervals span the region of  $\sin^2(\theta_{13}) = [0.008, 0.08]$ . This  
 3044 is more than an order of magnitude worse than the constraint from reactor  
 3045 experiments [75].

3046 As previously discussed, the correlations between oscillation parameters are  
 3047 also important to understand how the atmospheric samples respond. Figure 8.10  
 3048 illustrates the two dimensional  $\sin^2(\theta_{13}) - \delta_{CP}$  sensitivity, marginalised over all  
 3049 other parameters. The displayed contours are calculated by marginalising over  
 3050 both hierarchies. The shape of the  $1\sigma$  credible interval shows that the constraining  
 3051 power of the fit on  $\delta_{CP}$  is dependent upon the value of  $\sin^2(\theta_{13})$ . Furthermore,  
 3052 they show a strong resemblance to the likelihood scans illustrated in Figure 8.2.  
 3053 Whilst the atmospheric samples do not strongly constrain the value of  $\sin^2(\theta_{13})$ ,  
 3054 the value of  $\sin^2(\theta_{13})$  does impact the atmospheric sensitivity to  $\delta_{CP}$ . A value of  
 3055  $\sin^2(\theta_{13}) \sim 0.02$  would select a continuous contour over all values of  $\delta_{CP}$ . This  
 3056 shows the effect of the marginalisation effect previously described.

### Without Reactor Constraint, Both Hierarchies



**Figure 8.9:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{13})$ , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3057     The  $\sin^2(\theta_{23}) - \Delta m_{32}^2$  disappearance contours are illustrated in Figure 8.11.  
 3058     As expected, the area contained in the inverted hierarchy  $1\sigma$  credible interval is  
 3059     slightly smaller than that in the normal hierarchy. This follows from the Bayes  
 3060     factor showing a weak preference for NH meaning that more of the steps will exist  
 3061     in the  $\Delta m_{32}^2 > 0$  region. The known oscillation parameters of  $\sin^2(\theta_{23}) = 0.528$   
 3062     and  $\Delta m_{32}^2 = 2.509 \times 10^{-3} \text{ eV}^2$  are contained within the  $1\sigma$  credible interval.

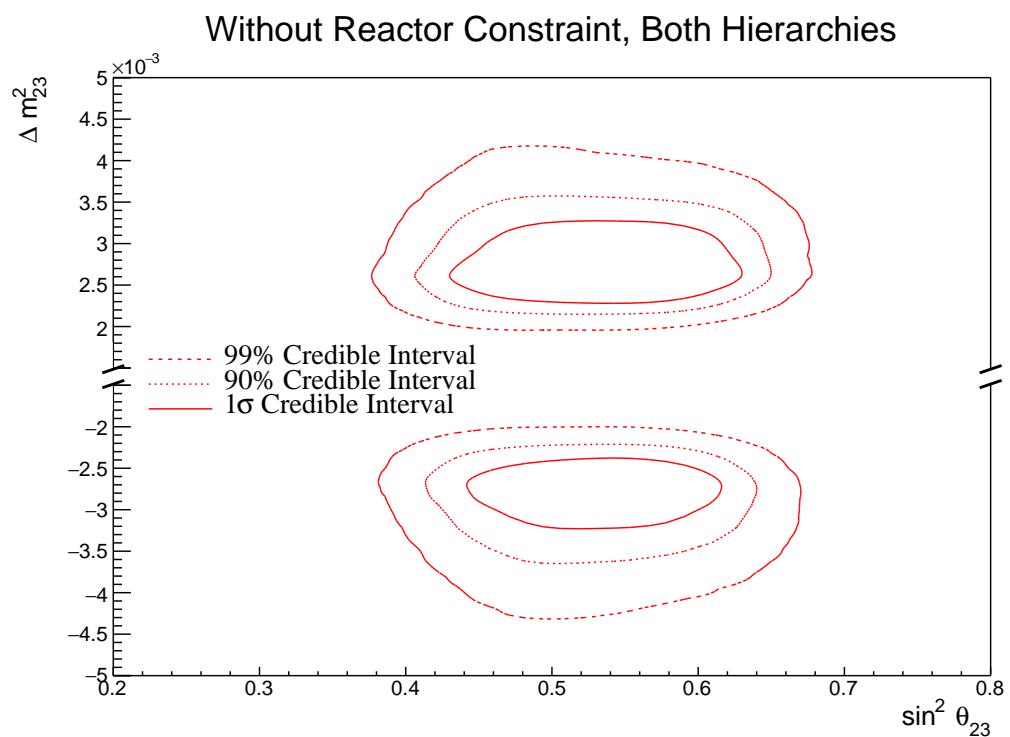
3063     Figure 8.12 illustrates the two-dimensional projections for each permutation of  
 3064     oscillation parameters which this analysis is sensitive to:  $\delta_{CP}$ ,  $\sin^2(\theta_{13})$ ,  $\sin^2(\theta_{23})$ ,  
 3065     and  $\Delta m_{32}^2$ . The purpose of this plot is to illustrate the correlations between the  
 3066     oscillation parameters. The contours are calculated whilst marginalising over  
 3067     both hierarchies, however, only the NH is illustrated when plotting the  $\Delta m_{32}^2$   
 3068     parameter. As expected the correlations play a significant role in these sensitivity  
 3069     measurements, especially the choice of the  $\sin^2(\theta_{13})$  constraint. The application  
 3070     of reactor constraint would be expected to alter both the width and position of the

### Without Reactor Constraint, Both Hierarchies

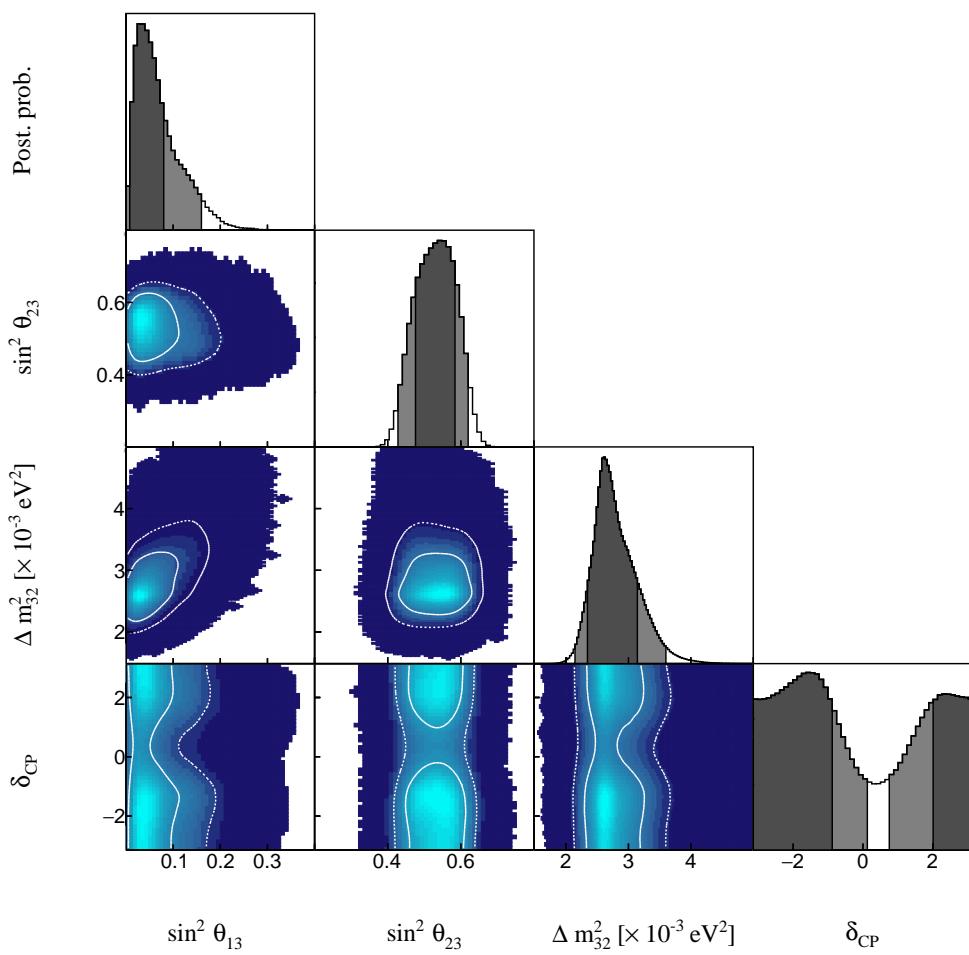


**Figure 8.10:** The two-dimensional posterior probability density distribution in  $\delta_{CP}$  –  $\sin^2(\theta_{13})$ , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.

3071  $\Delta m_{32}^2$ ,  $\delta_{CP}$ , and  $\sin^2(\theta_{23})$  constraints. The majority of the octant model preference  
 3072 comes from the region of  $\sin^2(\theta_{13}) \sim 0.03$  such that the application of the reactor  
 3073 constraint would not be expected to significantly change the octant preference.  
 3074 The reactor constraint would result in lower values of  $|\Delta m_{32}^2|$ . Interestingly, the  
 3075 distribution of steps in the  $\delta_{CP}$ - $\sin^2(\theta_{13})$  plot is slightly flatter in the region of the  
 3076 reactor constraint. Both the posterior distribution from this fit and the distribution  
 3077 in Figure 8.2 show a region of low negative log-likelihood extending out towards  
 3078 higher values of  $\sin^2(\theta_{13})$  in the  $\delta_{CP} \sim -\pi/2$  and  $\delta_{CP} \sim 2$  region. Consequently,  
 3079 the reactor constraint could feasibly reduce the sensitivity of the atmospheric  
 3080 samples to  $\delta_{CP}$ , due to the previously discussed marginalisation effects.



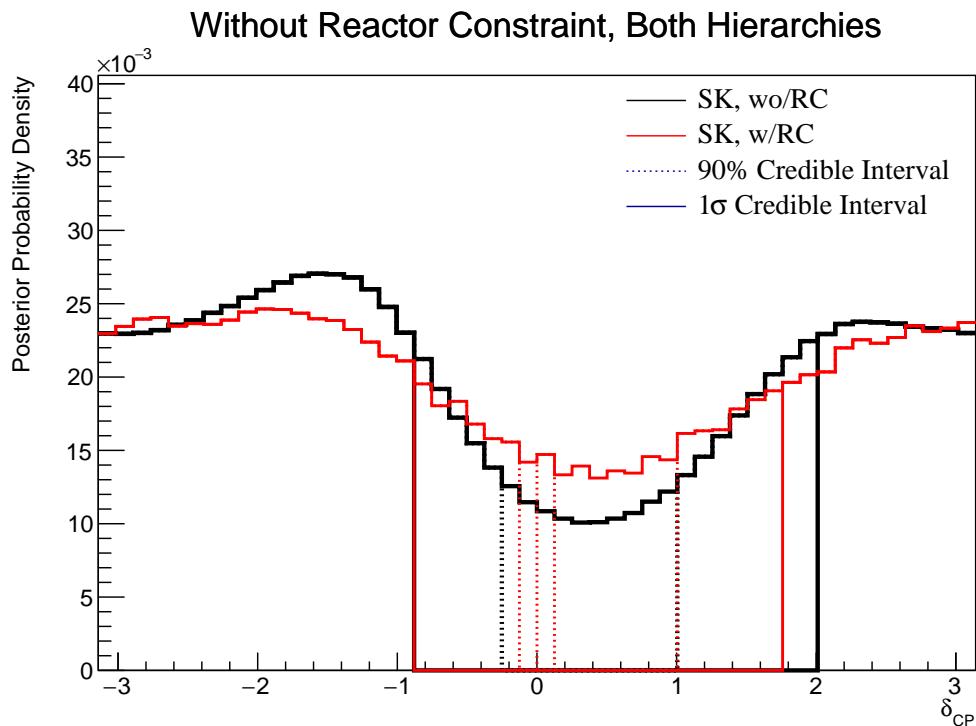
**Figure 8.11:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \sin^2(\theta_{23})$ , marginalised over both hierarchies, from the SK atmospheric-only fit. The reactor constraint is not applied.



**Figure 8.12:** The posterior probability density distribution from the SK atmospheric-only fit. The reactor constraint is not applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

### 3081 8.3.2 Atmospheric-Only Sensitivity With Reactor Constraint

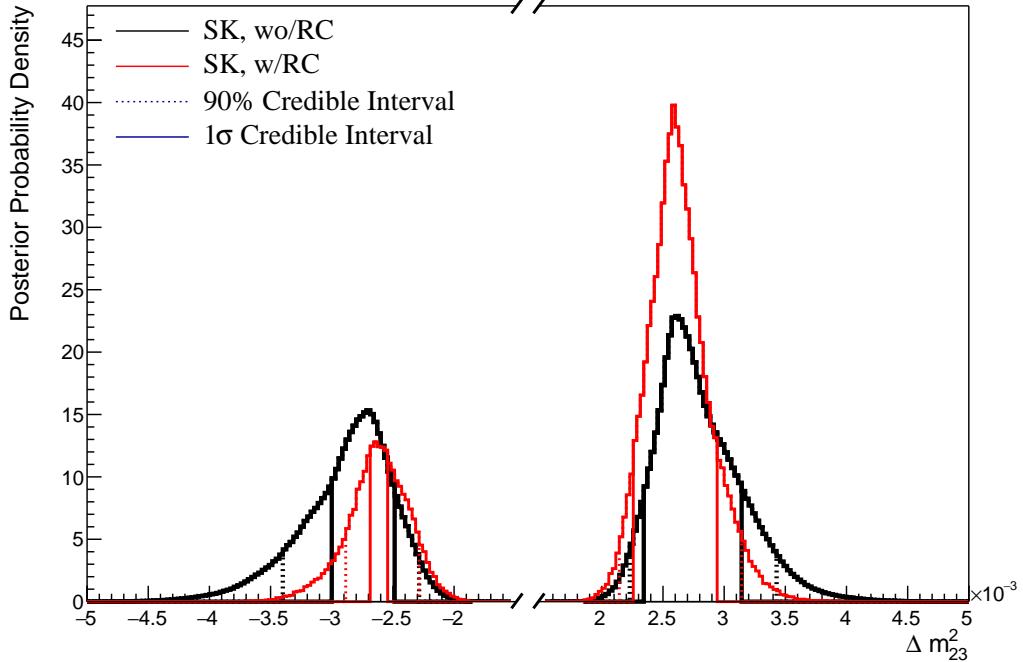
3082 The results in subsection 8.3.1 discuss the atmospheric sensitivity when the reactor  
 3083 constraint is not applied. The correlations illustrated in Figure 8.12 indicate that  
 3084 the marginalisation effects could contribute to differing sensitivities when the  
 3085 external reactor constraint is applied. Using the technique discussed in subsec-  
 3086 tion 4.1.1, the posterior distribution of the fit in subsection 8.3.1 can be reweighted  
 3087 to include the reactor constraint of  $\sin^2(\theta_{13}) = (2.18 \pm 0.08) \times 10^{-2}$  [75].



**Figure 8.13:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$  compared between the SK atmospheric-only fit (Black) and the SK atmospheric fit with the reactor constraint applied (Red). The distributions are marginalised over both hierarchies.

3088 Figure 8.13 illustrates the sensitivity to  $\delta_{CP}$  of the atmospheric fit with reactor  
 3089 constraint applied. The distribution is less peaked than the previous results.  
 3090 This is due to the expected marginalisation effect previously discussed. The  
 3091 width of the  $1\sigma$  credible interval is increased when the reactor constraint is  
 3092 applied, indicating less sensitivity to  $\delta_{CP}$  in the region of  $\sin^2(\theta_{13})$  preferred  
 3093 by the reactor constraint.

### Without Reactor Constraint, Both Hierarchies



**Figure 8.14:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  compared between the SK atmospheric-only fit (Black) and the SK atmospheric fit with the reactor constraint applied (Red). The distributions are marginalised over both hierarchies.

3094     The reactor constraint increases the sensitivity of the atmospheric samples to  
 3095      $\Delta m_{32}^2$  as illustrated in Figure 8.14. The  $1\sigma$  credible interval in  $\Delta m_{32}^2$  is determined  
 3096     to be  $[-2.70, -2.55] \times 10^{-3}$  eV<sup>2</sup> and  $[2.25, 2.95] \times 10^{-3}$  eV<sup>2</sup>. The width of the IH  
 3097     credible interval is reduced by  $\sim 70\%$  when the reactor constraint is applied. Due  
 3098     to the marginalisation effects observed in Figure 8.12, the favoured region of  $\Delta m_{32}^2$   
 3099     moves closer to zero for both hierarchies. A clear explanation of this behaviour  
 3100     is illustrated in Figure 8.15 which illustrates the posterior distribution in the  
 3101      $\Delta m_{32}^2 - \sin^2(\theta_{13})$  parameters, marginalised over both hierarchies. The correlation  
 3102     between  $\Delta m_{32}^2$  and  $\sin^2(\theta_{13})$  is such that lower values of  $\sin^2(\theta_{13})$  tend towards  
 3103     lower values of  $|\Delta m_{32}^2|$ . This moves the posterior distribution towards the known  
 3104     oscillation parameter  $\Delta m_{32}^2 = 2.509 \times 10^{-3}$  eV<sup>2</sup>.

3105     Table 8.4 presents the fraction of steps in each hierarchy and octant model  
 3106     for the fit after the reactor constraint has been applied. The reactor constraint  
 3107     significantly increases the NH preference, increasing the Bayes factor from



**Figure 8.15:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \sin^2(\theta_{13})$  compared between the SK atmospheric-only fit (Black) and the SK atmospheric fit with the reactor constraint (Red). The distributions are marginalised over both hierarchies.

	LO ( $\sin^2 \theta_{23} < 0.5$ )	UO ( $\sin^2 \theta_{23} > 0.5$ )	Sum
NH ( $\Delta m_{32}^2 > 0$ )	0.21	0.53	0.74
IH ( $\Delta m_{32}^2 < 0$ )	0.08	0.18	0.26
Sum	0.29	0.71	1.00

**Table 8.4:** The distribution of steps in an SK atmospheric with reactor constraint fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The Bayes factors are calculated as  $B(\text{NH}/\text{IH}) = 2.86$  and  $B(\text{UO}/\text{LO}) = 2.39$ .

<sup>3108</sup>  $B(\text{NH}/\text{IH}) = 1.37$  to  $B(\text{NH}/\text{IH}) = 2.86$  when the reactor constraint is applied.

<sup>3109</sup> This is still defined as a weak preference for NH hypothesis according to Jeffrey's

<sup>3110</sup> scale (see Table 4.1), however, it is a stronger preference than when the constraint

<sup>3111</sup> is not applied. The preference for the correct octant model is slightly increased

<sup>3112</sup> by the application of the reactor constraint which is consistent with expectation.

<sup>3113</sup> However, the conclusion that would be made does not significantly change.

### 3114    8.3.3 Application of Near Detector Constraints for Atmospheric 3115    Samples

3116    The choice of applying the near detector constraints to the low-energy atmo-  
3117    spheric samples was introduced in subsection 6.4.3. This subsection illustrates  
3118    the effect of that choice on the sensitivities of the atmospheric samples to the  
3119    oscillation parameters. This Asimov data was generated assuming the ‘AsimovA’  
3120    oscillation parameter set defined in Table 2.2 and the post-BANFF systematic  
3121    parameter tune.

3122    The change in sensitivity on  $\delta_{CP}$  is given in Figure 8.16. The reactor constraint  
3123    is not applied in either of the fits within this comparison. The shape of the  
3124    posterior is similar although less peaked at the Asimov point ( $\delta_{CP} = -1.601$ )  
3125    and more symmetric between the regions of  $\delta_{CP} = -1.601$  and  $\delta_{CP} \sim 2.5$ . The  
3126    width of the  $1\sigma$  credible intervals are approximately the same (identical to within  
3127    a bin width) and the same conclusion holds for the higher credible intervals. The  
3128    change in sensitivity to other oscillation parameters has been studied and no  
3129    significant discrepancies were found. As expected, the sensitivities are statistics  
3130    dominated and the exact choice of systematic model and constraint does not  
3131    significantly affect the physics conclusions one would make from this analysis.

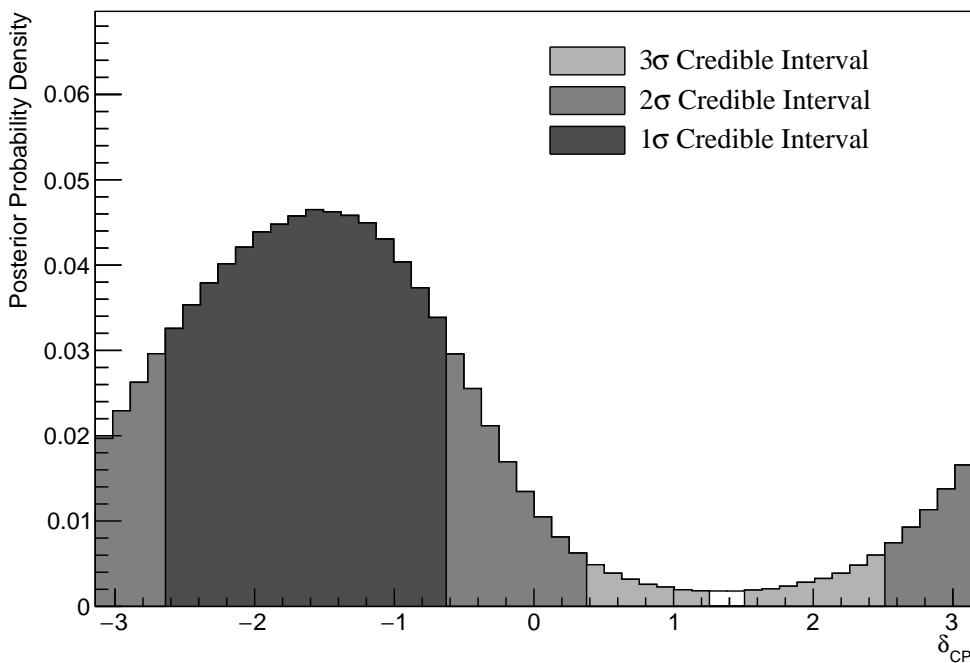


**Figure 8.16:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$  compared between the SK atmospheric-only fit where the near detector constraint is (Black) and is not (Red) applied. The distributions are marginalised over both hierarchies.

### 3132 8.3.4 Atmospheric and Beam Sensitivity without Reactor Con- 3133 straint

3134 This section presents the sensitivities of the simultaneous beam and atmospheric  
 3135 analysis where the reactor constraint is not applied. Similar to the previous  
 3136 studies, the Asimov data is built assuming the post-BANFF cross-section tune  
 3137 and Asimov A oscillation parameters defined in Table 2.2. This fit uses all 18 near  
 3138 detector beam samples, 5 far detector beam samples, and 18 atmospheric samples.  
 3139 The sensitivity to  $\delta_{CP}$ , marginalised over both hierarchies, is given in Figure 8.17.

Without Reactor Constraint, Both Hierarchies



**Figure 8.17:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$ , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

3140 The credible intervals and highest posterior distribution for each oscillation  
 3141 parameter is given in Table 8.5. The highest posterior probability density is  
 3142  $\delta_{CP} = -1.58$  and is compatible with the known Asimov A value of  $\delta_{CP} = -1.601$ .  
 3143 The CP-conserving values of  $\delta_{CP} = 0, \pi, -\pi$  are disfavoured at  $1\sigma$  credible  
 3144 interval. There is also a region around  $\delta_{CP} = 1.4$  which is disfavoured at more  
 3145 than  $3\sigma$ . Whilst these conclusions can only be made at this particular Asimov

3146 point, it does show that if the true value of  $\delta_{CP}$  was CP-violating, this joint  
3147 analysis would be able to disfavour CP conserving values at over  $1\sigma$  without  
3148 any external constraints. The highest posterior probability density does move  
3149 further away from the Asimov point when only steps in the NH region are  
3150 considered. This is due to the correlations between the value of  $\delta_{CP}$  and the  
3151 mass hierarchy, as will be later discussed.

Parameter	Interval	HPD
$\delta_{CP}$ , (BH)	$[-2.64, -0.63]$	-1.57
$\delta_{CP}$ , (NH)	$[-2.76, -0.63]$	-1.45
$\delta_{CP}$ , (IH)	$[-2.39, -0.88]$	-1.57
$\Delta m_{32}^2$ (BH) [ $\times 10^{-3}\text{eV}^2$ ]	[2.46, 2.58]	2.49
$\Delta m_{32}^2$ (NH) [ $\times 10^{-3}\text{eV}^2$ ]	[2.48, 2.56]	2.51
$\Delta m_{32}^2$ (IH) [ $\times 10^{-3}\text{eV}^2$ ]	$[-2.60, -2.52]$	-2.55
$\sin^2(\theta_{23})$ (BH)	[0.48, 0.55]	0.509
$\sin^2(\theta_{23})$ (NH)	[0.48, 0.55]	0.509
$\sin^2(\theta_{23})$ (IH)	[0.48, 0.55]	0.521

**Table 8.5:** The position of the highest posterior probability density (HPD) and width of the  $1\sigma$  credible interval for the joint beam and atmospheric fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

3152 The sensitivity to  $\Delta m_{32}^2$  is illustrated in Figure 8.18, marginalised over both  
3153 hierarchies. Notably, the  $1\sigma$  credible interval is entirely contained within the  
3154 normal hierarchy region, as illustrated in Table 8.5. This is illustrates reasonable  
3155 sensitivity to the mass hierarchy model. This is also reflected in the  $1\sigma$  credible  
3156 intervals being approximately the same when they are made considering both  
3157 hierarchies and when considering only the NH. The known oscillaton parameter  
3158 is  $\Delta m_{32}^2 = 2.509 \times 10^{-3}\text{eV}^2$ . The normal hierarchy distribution favours this value  
3159 with the highest posterior probability density of  $\Delta m_{32}^2 = 2.51 \times 10^{-3}\text{eV}^2$ .

3160 The fraction of steps in each of the mass hierarchy regions and octants of  
3161  $\sin^2(\theta_{23})$  is given in Table 8.6. The Bayes factors are determined to be  $B(\text{NH}/\text{IH}) =$

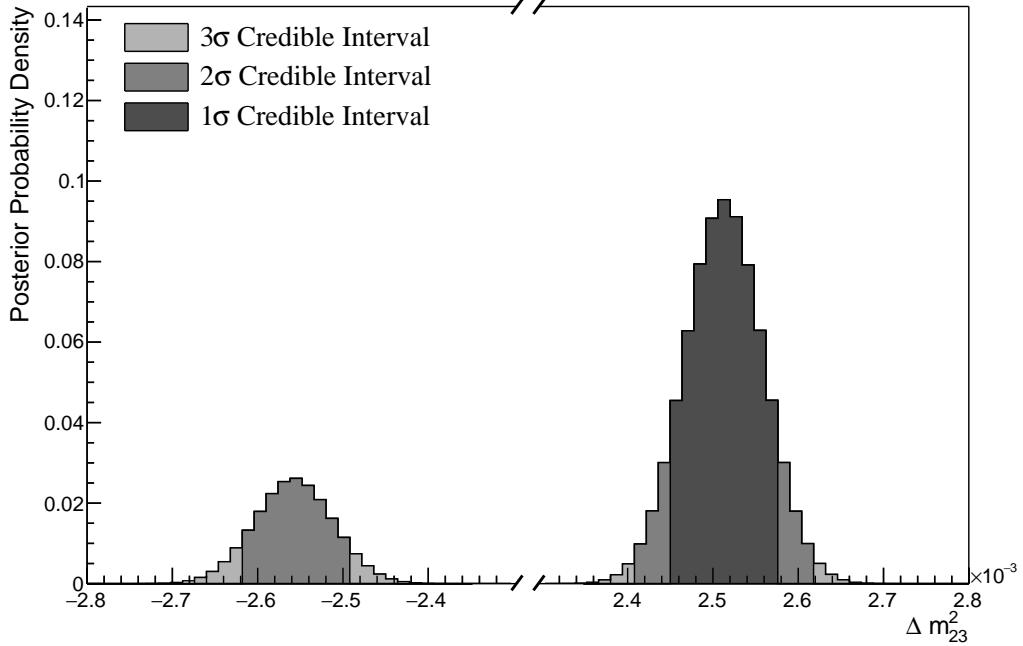
3162 3.67 and  $B(\text{UO}/\text{LO}) = 1.74$ . Jeffrey's scale (presented in Table 4.1) states that  
 3163 this value of the hierarchy Bayes factor illustrates substantial evidence for the  
 3164 normal hierarchy hypothesis. This corresponds to the correct hypothesis given  
 3165 the known oscillation parameters. It is a stronger statement than the atmospheric-  
 3166 only analysis can provide. It is important to note that this is a substantial  
 3167 preference that requires no external constraints required. The Bayes factor for  
 3168 octant determination represents a weak preference for the upper octant but does  
 3169 select the correct octant model.

	LO ( $\sin^2 \theta_{23} < 0.5$ )	UO ( $\sin^2 \theta_{23} > 0.5$ )	Sum
NH ( $\Delta m_{32}^2 > 0$ )	0.29	0.50	0.79
IH ( $\Delta m_{32}^2 < 0$ )	0.08	0.13	0.21
Sum	0.37	0.63	1.00

**Table 8.6:** The distribution of steps in a joint beam and atmospheric fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as  $B(\text{NH}/\text{IH}) = 3.67$  and  $B(\text{UO}/\text{LO}) = 1.74$ .

3170 The sensitivity to  $\sin^2(\theta_{23})$  is presented in Figure 8.19. There is a clear  
 3171 preference for the upper octant but the peak of the distribution is relatively  
 3172 flat. It peaks at  $\sin^2(\theta_{23}) = 0.509$  which is in the region of the known value of  
 3173  $\sin^2(\theta_{23}) = 0.528$ . The difference in the highest posterior distribution and the  
 3174 width of the credible interval is relatively unchanged when considering different  
 3175 hierarchy models showing no strong correlation between  $\sin^2(\theta_{23})$  and  $|\Delta m_{32}^2|$ .  
 3176 The sensitivity presented as a function of the appearance parameters ( $\sin^2(\theta_{13}) -$   
 3177  $\delta_{CP}$ ) is given in Figure 8.20. As expected, the contours follow that given in  
 3178 Figure 8.2, where the  $2\sigma$  credible intervals have a closed contour excluding the  
 3179 region around  $\delta_{CP} \sim 1.2$ . The width of the  $3\sigma$  credible interval is also clearly  
 3180 dependent upon the value of  $\delta_{CP}$ . Close to the Asimov point,  $\delta_{CP} = -1.601$ , the  
 3181 width of the  $3\sigma$  credible interval approximately spans  $\sin^2(\theta_{13}) = [0.013, 0.04]$ .  
 3182 This is reduced to a region of  $\sin^2(\theta_{13}) = [0.023, 0.042]$  at the most disfavoured  
 3183 value of  $\delta_{CP}$ . This follows the behaviour shown in the likelihood scans. The  $1\sigma$

### Without Reactor Constraint, Both Hierarchies



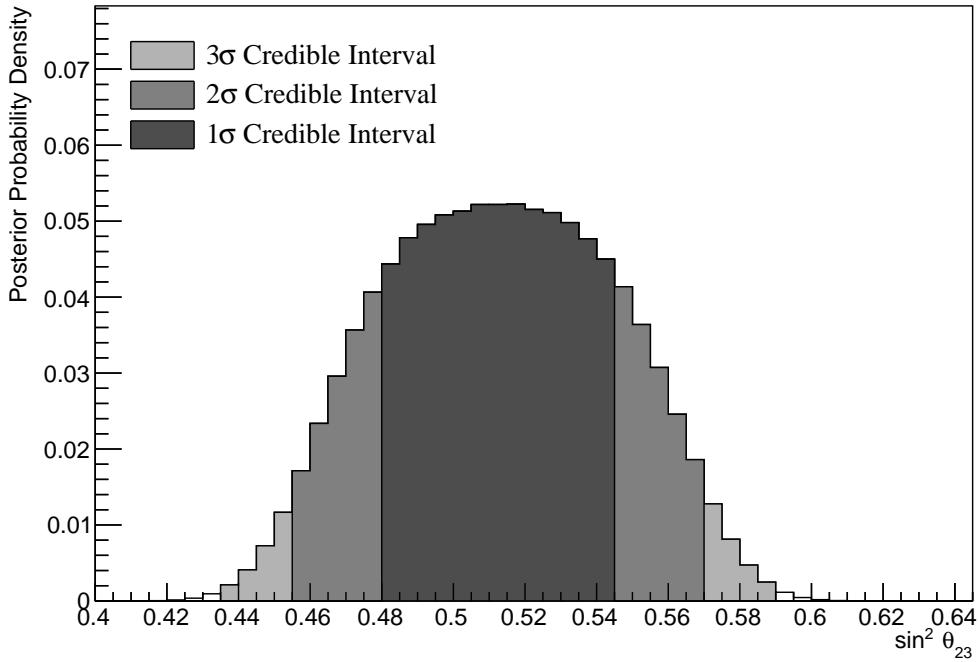
**Figure 8.18:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$ , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

credible interval is consistent with both the known oscillation parameter and the reactor constraint ( $\sin^2(\theta_{13}) = (2.18 \pm 0.08) \times 10^{-2}$ ). Application of the reactor constraint would be expected to decrease the width of the 1 $\sigma$  credible intervals of  $\delta_{CP}$  due to the triangular shape of the posterior probability.

The sensitivity in terms of the ‘disappearance’ parameters marginalised over both hierarchies is given in Figure 8.21. The area contained within the IH credible intervals is significantly smaller than those in the NH region. This is reflected in the IH credible intervals being tighter in the  $\sin^2(\theta_{23})$  dimension. No significant correlation is observed between the value of  $\sin^2(\theta_{23})$  and  $|\Delta m_{32}^2|$ .

The two-dimensional posterior distribution for each permutation of the oscillation parameters of interest is given in Figure 8.22. The most notable observation is that the  $\sin^2(\theta_{13})$  and  $\sin^2(\theta_{23})$  are anti-correlated. If the value of  $\sin^2(\theta_{13})$  was known to be closer to the known oscillation parameter value, the preferred value of  $\sin^2(\theta_{23})$  would increase furthering the preference for the UO. That

### Without Reactor Constraint, Both Hierarchies



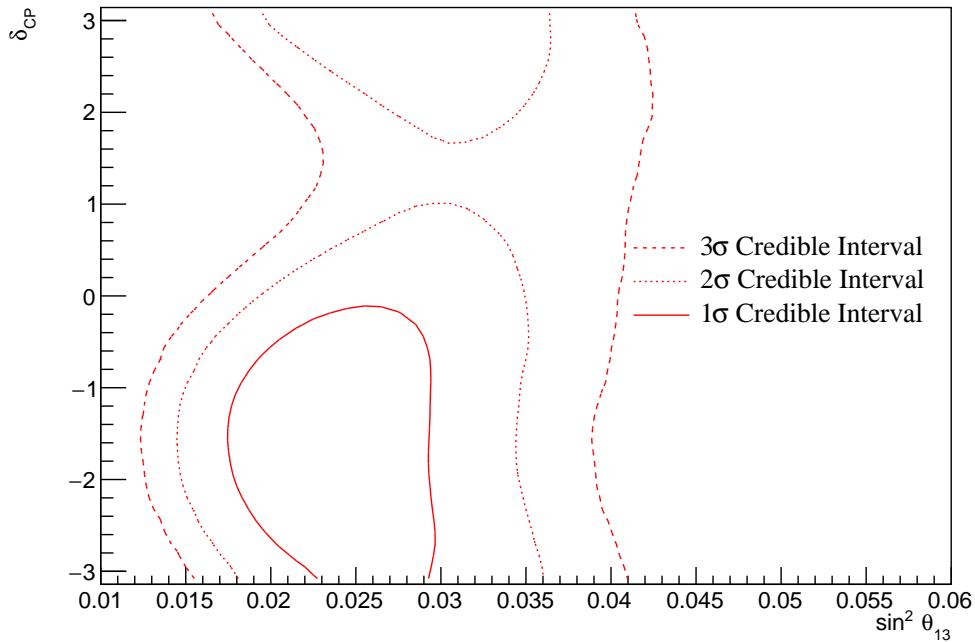
**Figure 8.19:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{23})$ , marginalised over both hierarchies, from the joint beam-atmospheric fit. The reactor constraint is not applied.

3198 would move the highest posterior probability closer in line with the Asimov  
 3199 value. This also means that the preference for the UO would be increased if  
 3200 the reactor constraint was to be applied.

3201 Furthermore, the  $\delta_{CP}$  and  $|\Delta m_{32}^2|$  oscillation parameters are anti-correlated,  
 3202 such that higher values of  $|\Delta m_{32}^2|$  prefer lower values of  $\delta_{CP}$ . Whilst this is  
 3203 an interesting result on its own, the width of the  $\Delta m_{32}^2$  contours also depend  
 3204 on  $\sin^2(\theta_{13})$ . This introduces another correlation effect that could modify the  
 3205 sensitivity to  $\delta_{CP}$  once the reactor constraint is applied.

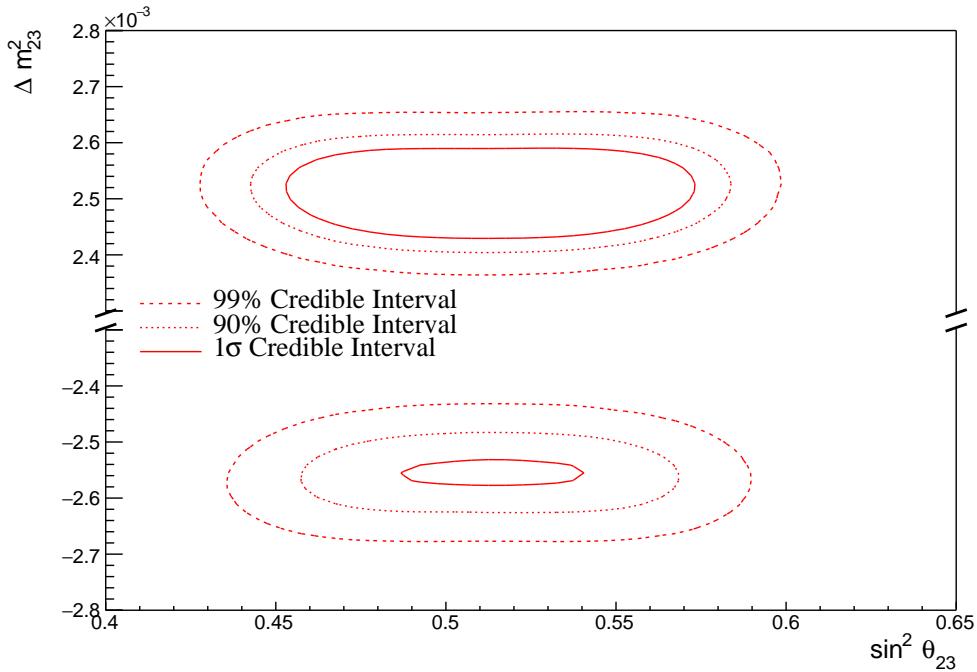
3206 The correlation between  $\sin^2(\theta_{13})$  and  $\Delta m_{32}^2$  can be seen in Figure 8.23. A  
 3207 much larger fraction of the posterior distribution is contained in the NH for lower  
 3208 values of  $\sin^2(\theta_{13})$ . Consequently, the application of the reactor constraint would  
 3209 be expected to significantly increase the preference for NH.

### Without Reactor Constraint, Both Hierarchies

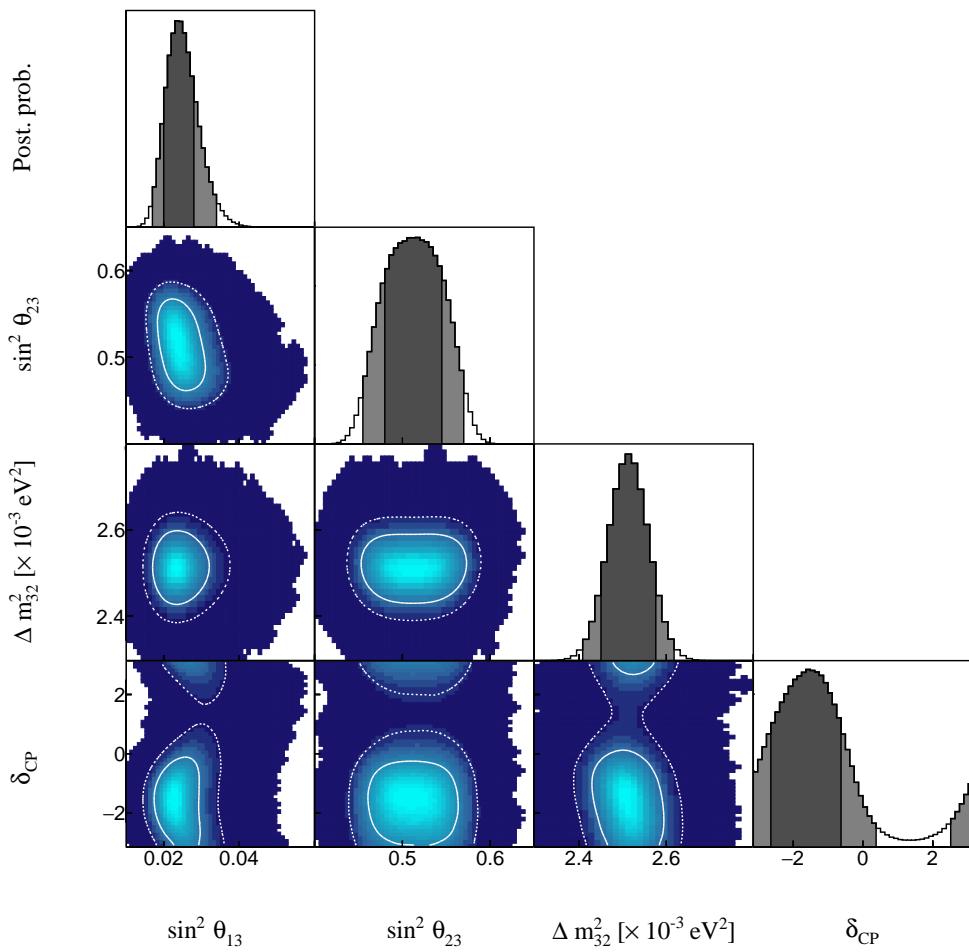


**Figure 8.20:** The two-dimensional posterior probability density distribution in  $\delta_{CP}$  –  $\sin^2(\theta_{13})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

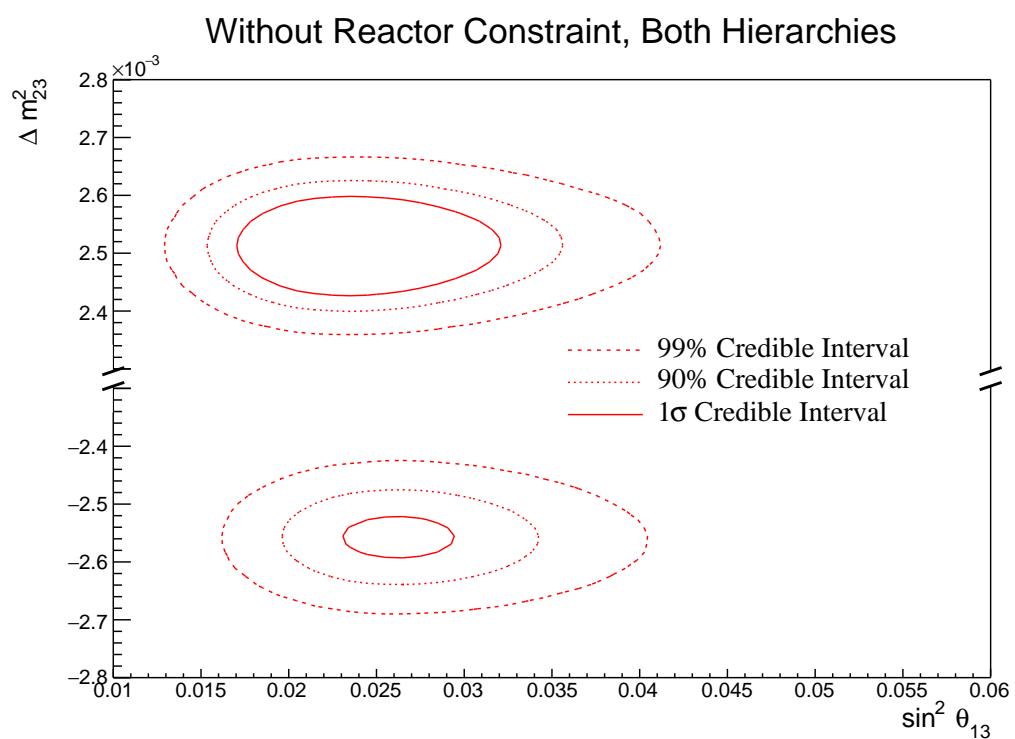
### Without Reactor Constraint, Both Hierarchies



**Figure 8.21:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  –  $\sin^2(\theta_{23})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.



**Figure 8.22:** The posterior probability density distribution from the joint beam and atmospheric fit. The reactor constraint is not applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

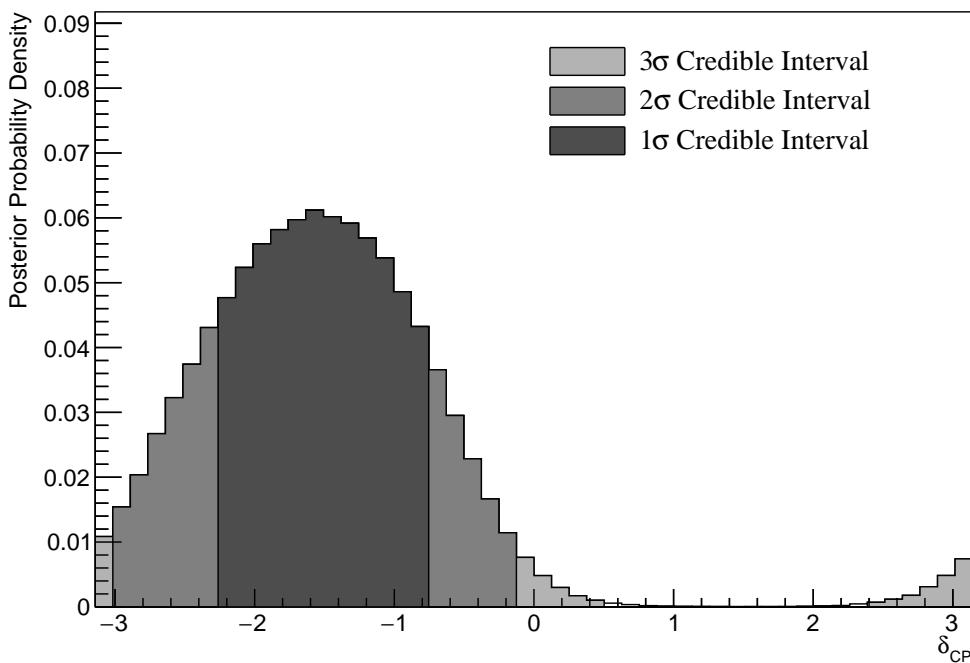


**Figure 8.23:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \sin^2(\theta_{13})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit. The reactor constraint is not applied.

### 3210 8.3.5 Atmospheric and Beam Sensitivity with Reactor Constraint

3211 This section presents the sensitivities of the joint beam and atmospheric fit when  
 3212 the reactor constraint is applied to  $\sin^2(\theta_{13})$ . As with the previous studies, the  
 3213 Asimov data is made using the AsimovA oscillation parameter set defined in  
 3214 Table 2.2 and the post-BANFF systematic parameter tune.

With Reactor Constraint, Both Hierarchies

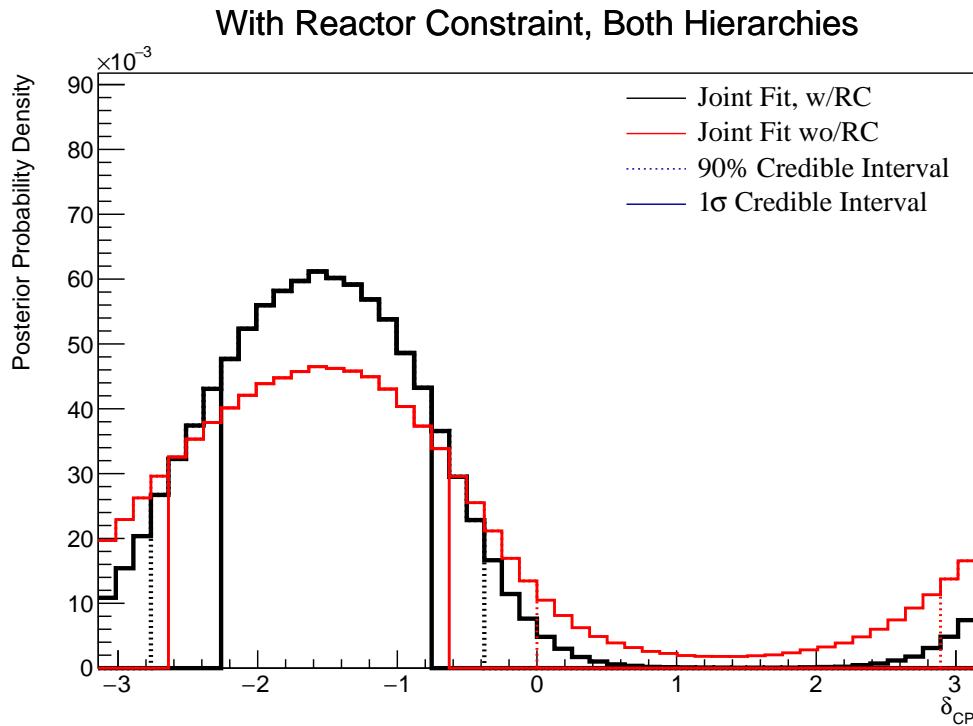


**Figure 8.24:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$ , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

3215 Figure 8.24 illustrates the sensitivity to  $\delta_{CP}$ , marginalised over both hierarchies.  
 3216 The CP-conserving values of  $\delta_{CP} = -\pi, 0, \pi$  are disfavoured at  $2\sigma$ . Furthermore,  
 3217 the  $3\sigma$  credible interval excludes the region of  $\delta_{CP} = [0.50, 2.39]$ . Thus clearly  
 3218 disfavouring the region of  $\delta_{CP} = \pi/2$  at more than  $3\sigma$  for this particular set  
 3219 of known oscillation parameters. The width of the  $1\sigma$  credible intervals and  
 3220 the position of the highest posterior probability density is given in Table 8.7.  
 3221 The highest posterior probability density in  $\delta_{CP}$  is calculated as  $\delta_{CP} = -1.57$   
 3222 showing no significant biases in the determination of the known oscillation

parameters. The posterior distribution is more peaked around the known oscillation parameter value of  $\delta_{CP} = -1.601$ , as compared to the sensitivities when the reactor constraint is not applied (subsection 8.3.4). This follows from the correlations shown in Figure 8.20, where a lower value of  $\sin^2(\theta_{13})$  results in tighter constraints on  $\delta_{CP}$ .

The effect of applying the reactor constraint for  $\delta_{CP}$  in the joint beam-atmospheric fit is presented in Figure 8.25. The posterior distribution from the two fits are marginalised over both hierarchies. Clearly, the reactor constraint improves the ability of the fit to select the known oscillation parameter as the shape of the distribution is much more peaked. This is also evidenced by the tightening of the  $1\sigma$  and 90% credible intervals. Additionally, the disfavoured region of  $1 < \delta_{CP} < 2$  is wider when the reactor constraint is applied.



**Figure 8.25:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$  compared between the joint beam-atmospheric fit (Red) and the joint beam and atmospheric fit with the reactor constraint (Black). The distributions are marginalised over both hierarchies.

The sensitivity to  $\sin^2(\theta_{23})$ , marginalised over both hierarchies, is given in Figure 8.26. The highest posterior probability density is located at  $\sin^2(\theta_{23}) = 0.527$

Parameter	Interval	HPD
$\delta_{CP}$ , (BH)	[-2.26, -0.75]	-1.57
$\delta_{CP}$ , (NH)	[-2.26, -0.75]	-1.57
$\delta_{CP}$ , (IH)	[-2.13, -1.00]	-1.57
$\Delta m_{32}^2$ (BH) [ $\times 10^{-3}\text{eV}^2$ ]	[2.46, 2.52]	2.49
$\Delta m_{32}^2$ (NH) [ $\times 10^{-3}\text{eV}^2$ ]	[2.48, 2.56]	2.51
$\Delta m_{32}^2$ (IH) [ $\times 10^{-3}\text{eV}^2$ ]	[-2.60, -2.52]	-2.55
$\sin^2(\theta_{23})$ (BH)	[0.49, 0.55]	0.527
$\sin^2(\theta_{23})$ (NH)	[0.49, 0.55]	0.527
$\sin^2(\theta_{23})$ (IH)	[0.50, 0.56]	0.539

**Table 8.7:** The position of the highest posterior probability density (HPD) and width of the  $1\sigma$  credible interval for the joint beam and atmospheric fit where the reactor constraint is applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH), and inverted hierarchy only (IH).

which agrees with the known value of  $\sin^2(\theta_{23}) = 0.528$ . The distribution clearly favours the UO with almost the entirety of the  $1\sigma$  credible interval contained in the region. Figure 8.27 highlights the sensitivity of the joint fit both with and without the reactor constraint. The fit where the reactor constraint is applied selects the known value much better ( $\sin^2(\theta_{23}) = 0.528$ ). Furthermore, the reactor constraint increases the UO preference which is evidenced by the distribution moving further away from the octant boundary. This indicates that there are marginalisation effects between the two mixing parameters. This follows from the correlation illustrated between  $\sin^2(\theta_{23}) - \sin^2(\theta_{13})$  in Figure 8.22. The posterior distribution of the fit with reactor constraint is more peaked compared to the flatter distribution when the reactor constraint is not applied.

The fraction of steps contained within the two hierarchy and two octant models is given in Table 8.8. The reactor constraint significantly reduces the fraction of steps that are contained within the IH-LO region from 0.08 to 0.02, whilst significantly increasing the fraction of steps within the NH-UO region from 0.53 to 0.64. The application of the reactor constraint increases the Bayes factor

### With Reactor Constraint, Both Hierarchies



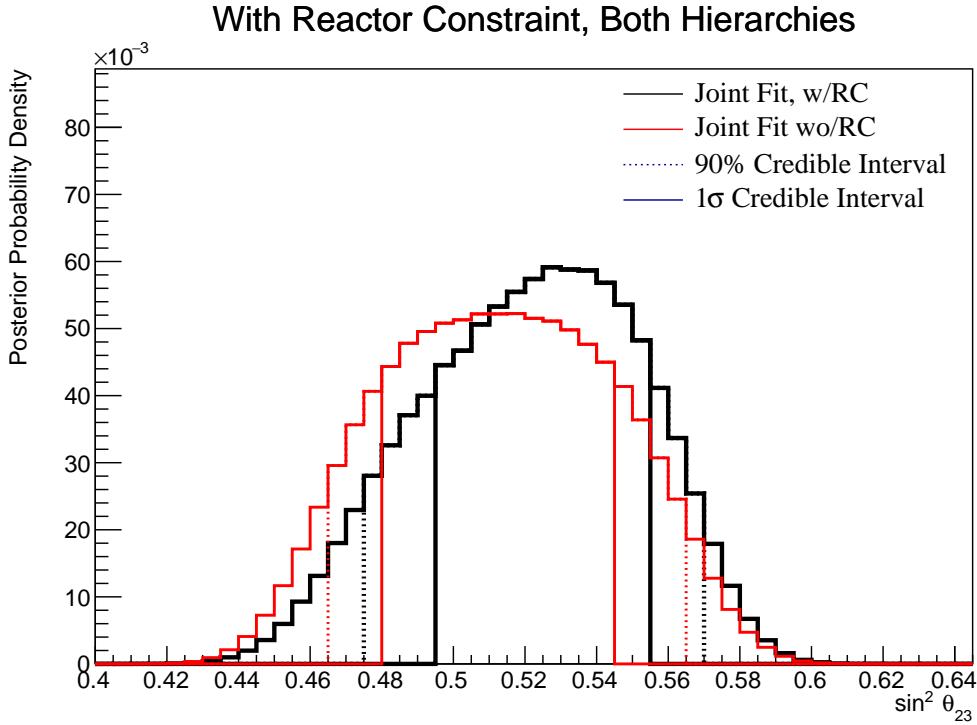
**Figure 8.26:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{23})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

from  $B(\text{NH}/\text{IH}) = 3.67$  to  $B(\text{NH}/\text{IH}) = 7.29$ . There is a very clear preference for the NH, with the Jeffreys scale stating a substantial preference for both fits (see subsection 4.3.3). The Bayes factor for UO preference is calculated as  $B(\text{UO}/\text{LO}) = 2.86$ . Whilst still a weak preference, this is certainly a stronger statement than the sensitivity when the reactor constraint is not applied.

	LO ( $\sin^2 \theta_{23} < 0.5$ )	UO ( $\sin^2 \theta_{23} > 0.5$ )	Sum
NH ( $\Delta m_{32}^2 > 0$ )	0.24	0.64	0.88
IH ( $\Delta m_{32}^2 < 0$ )	0.02	0.10	0.12
Sum	0.26	0.74	1.00

**Table 8.8:** The distribution of steps in a joint beam and atmospheric with the reactor constraint fit applied, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The Bayes factors are calculated as  $B(\text{NH}/\text{IH}) = 7.29$  and  $B(\text{UO}/\text{LO}) = 2.86$ .

The sensitivity to  $\Delta m_{32}^2$ , with the reactor constraint applied, is presented in



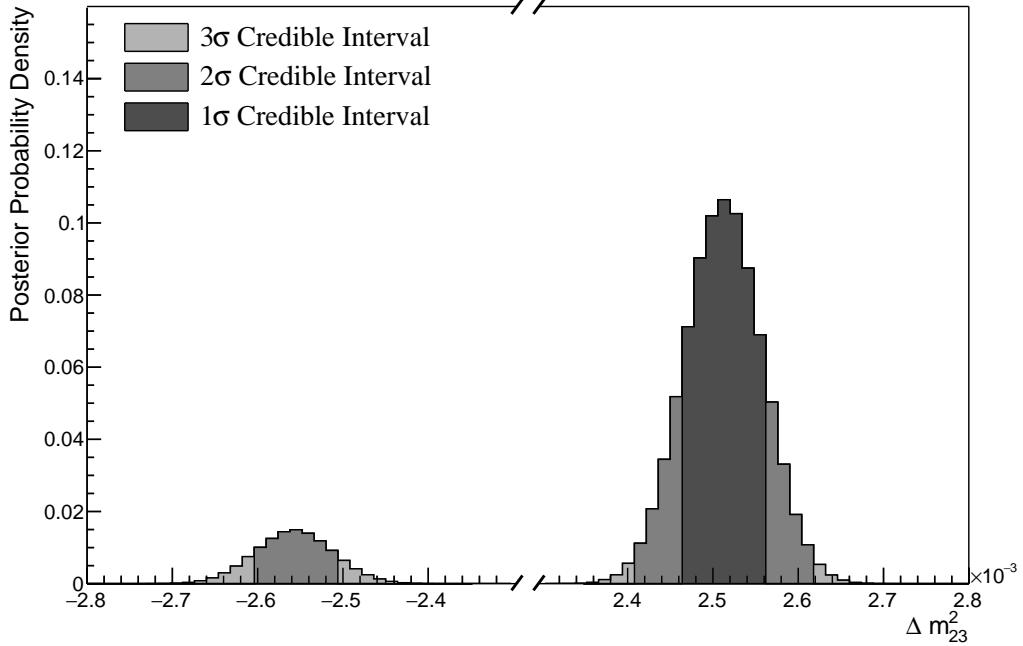
**Figure 8.27:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Red) and the joint beam and atmospheric fit with the reactor constraint (Black). The distributions are marginalised over both hierarchies.

3259 Figure 8.28. The posterior distribution is marginalised over both hierarchies. As  
 3260 expected, the  $1\sigma$  credible interval is entirely contained within the NH region. The  
 3261 position of the highest posterior probability density is given as  $2.49 \times 10^{-3} \text{ eV}^2$ ,  
 3262 illustrating no significant bias between the fit results and the known oscillation  
 3263 parameters. The application of the reactor constraint does move significantly the  
 3264 position of the credible intervals but does reduce their width.

3265 The sensitivity to the appearance parameters ( $\sin^2(\theta_{13}) - \delta_{CP}$ ) is given in  
 3266 Figure 8.29. The distribution is mostly uncorrelated between the two parameters  
 3267 and is centered at the known oscillation parameters. The  $1\sigma$  credible interval  
 3268 excludes  $\delta_{CP} = 0$  and  $\delta_{CP} = (-)\pi$ . Furthermore, the  $3\sigma$  credible intervals  
 3269 exclude the region of  $\delta_{CP} = \pi/2$ .

3270 The sensitivity to the disappearance parameters ( $\sin^2(\theta_{23}) - \Delta m_{32}^2$ ) is illus-  
 3271 trated in Figure 8.30. As expected from the one-dimensional distribution, the  $1\sigma$   
 3272 credible interval is entirely contained within the NH region. Both the NH and

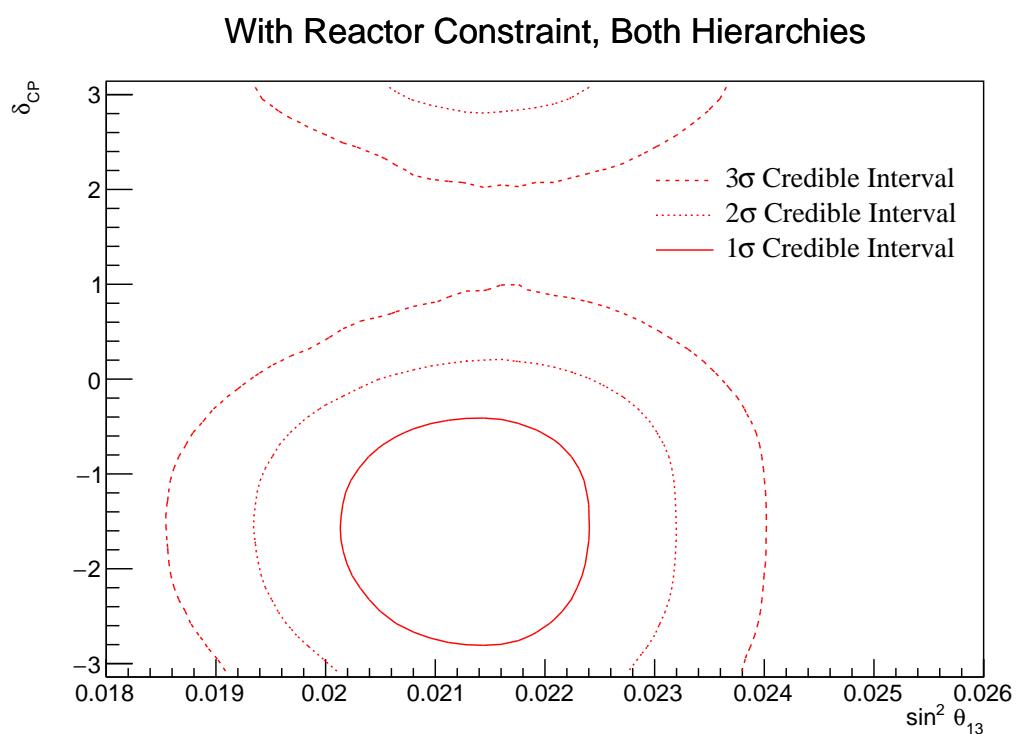
### With Reactor Constraint, Both Hierarchies



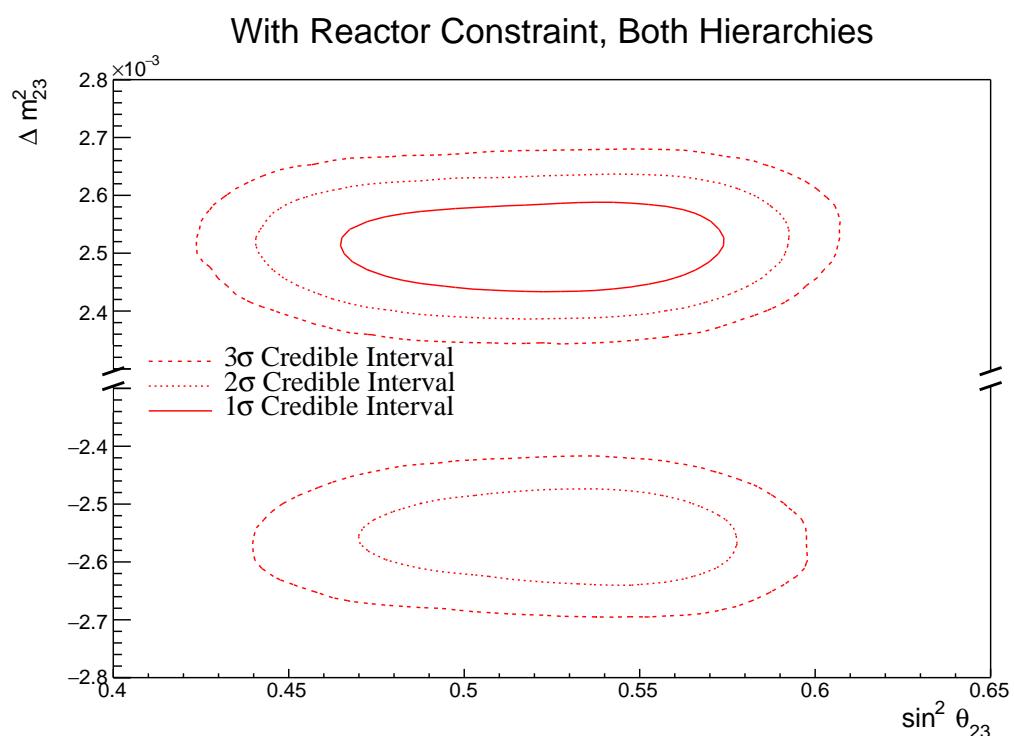
**Figure 8.28:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$ , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.

3273 IH regions favour the UO, with a visually similar preference in both hierarchies.  
 3274 The width of the  $\Delta m_{32}^2$  1 $\sigma$  credible interval does not significantly depend upon  
 3275 the value or octant of  $\sin^2(\theta_{23})$ . This shows that there are no strong correlations  
 3276 between these two parameters.

3277 Figure 8.31 illustrates the posterior distribution for each permutation of  
 3278 two oscillation parameters of interest. The application of the reactor constraint  
 3279 significantly reduces the correlations previously seen in Figure 8.22. There is  
 3280 still a small correlation between  $\delta_{CP}$  and  $\Delta m_{32}^2$ . The application of the reactor  
 3281 constraint has not significantly affected this correlation. The width of the 1 $\sigma$   
 3282 credible interval in  $\Delta m_{32}^2$  is wider for a value of  $\delta_{CP} = 0$  as compared to a value  
 3283 of  $\delta_{CP} = \pi$ . Similarly, the width of the 1 $\sigma$  credible interval in  $\delta_{CP}$  is smaller  
 3284 for lower values of  $\sin^2(\theta_{23})$ .



**Figure 8.29:** The two-dimensional posterior probability density distribution in  $\delta_{CP}$  –  $\sin^2(\theta_{13})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.



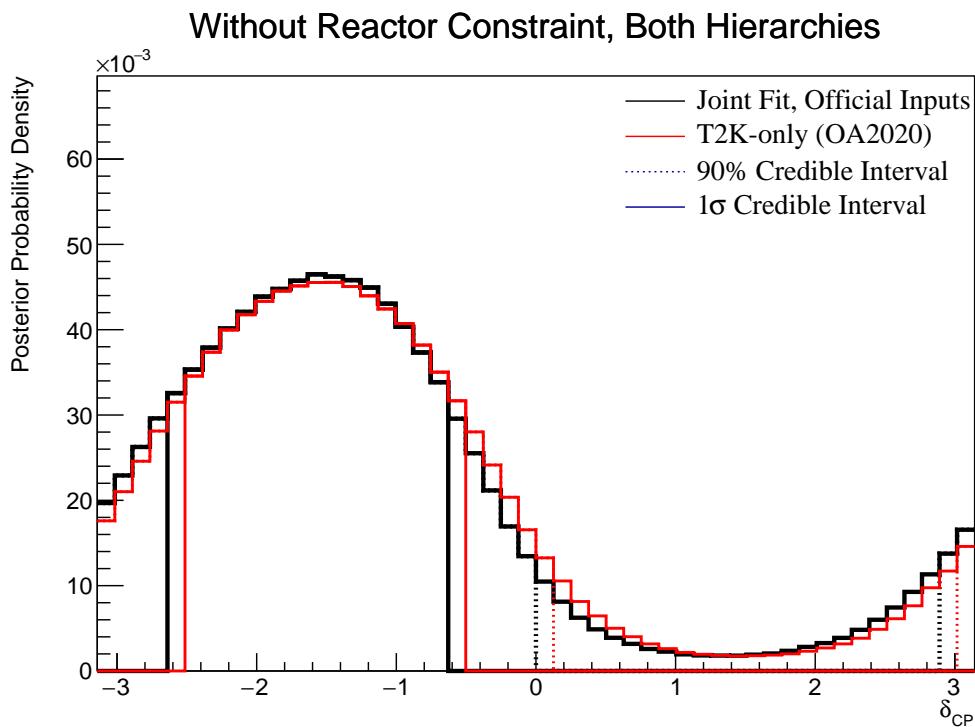
**Figure 8.30:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \sin^2(\theta_{23})$ , marginalised over both hierarchies, from the joint beam and atmospheric fit where the reactor constraint is applied.



**Figure 8.31:** The posterior probability density distribution from the joint beam and atmospheric fit where the reactor constraint is applied. The distribution is given for each two-dimensional permutation of the oscillation parameters of interest. The one-dimensional distribution of each parameter is also given.

### 3285 8.3.6 Comparison to Latest T2K Sensitivities without Reactor 3286 Constraint

3287 The benefits of the joint beam and atmospheric analysis can be determined by  
 3288 comparing the sensitivities to the beam-only analysis. This section presents those  
 3289 comparisons for sensitivities built using the Asimov A oscillation parameters  
 3290 defined in Table 2.2 and the post-BANFF systematic tune. The reactor constraint  
 3291 is not applied within either of the fits used in these comparisons.



**Figure 8.32:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

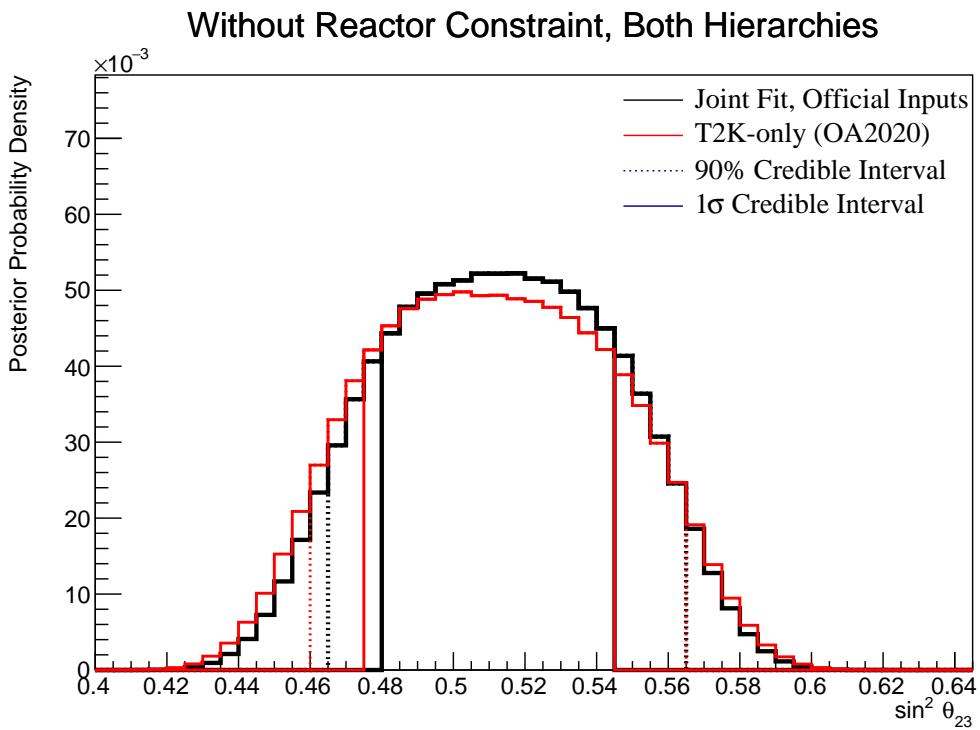
3292 The sensitivity, marginalised over both hierarchies, to  $\delta_{CP}$  from the joint  
 3293 beam-atmospheric and beam-only fits is presented in Figure 8.32. As expected  
 3294 from the likelihood scans (Figure 8.4), the sensitivity to  $\delta_{CP}$  is not significantly  
 3295 increased. This is because the known oscillation parameter value ( $\delta_{CP} = -1.601$ )  
 3296 lies at the position where the beam samples dominate the sensitivity compared  
 3297 to the SK samples.

The sensitivity to  $\Delta m_{32}^2$  of the joint beam-atmospheric fit is illustrated in Figure 8.33, where the posterior distribution has been marginalised over both hierarchies. The  $1\sigma$  credible interval of the joint beam and atmospheric fit is entirely contained within the NH region. This shows the significant increase in the ability of the fit to determine the correct mass hierarchy, as compared to the beam-only analysis. This is further evidenced by the fact that the 90% credible intervals from the joint fit are also tighter in the IH region as compared to the beam-only analysis. The Bayes factor for mass hierarchy determination for the beam-only and joint beam and atmospheric are  $B(\text{NH}/\text{IH}) = 1.91$  and  $B(\text{NH}/\text{IH}) = 3.67$ , respectively. According to Jeffrey's scale (Table 4.1), the beam-only analysis represents a weak preference for the NH hypothesis whereas the joint fit returns a substantial preference for the NH hypothesis. To summarise, the joint beam-atmospheric fit has a substantial preference for the correct hierarchy without the requirement of external constraints.



**Figure 8.33:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3312 The sensitivity to  $\sin^2(\theta_{23})$ , marginalised over both hierarchies, for both the  
 3313 beam-only and joint beam and atmospheric analysis are presented in Figure 8.34.  
 3314 The peak of the posterior distribution from the joint analysis is more aligned  
 3315 with the known value of  $\sin^2(\theta_{23}) = 0.528$  as compared to the beam-only  
 3316 analysis. This indicates that the marginalisation effects from other oscillation  
 3317 parameters ( $\sin^2(\theta_{13}) - \sin^2(\theta_{23})$ ) presented in Figure 8.22) are less prevalent in  
 3318 the projection of this parameter. The Bayes factors for the beam-only and joint  
 3319 beam-atmospheric fit are  $B(\text{UO}/\text{LO}) = 1.56$  and  $B(\text{UO}/\text{LO}) = 1.74$ , respectively.  
 3320 Therefore, the joint beam-atmospheric fit does prefer the UO more strongly than  
 3321 the beam-only analysis, albeit slightly.



**Figure 8.34:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3322 Whilst the beam-only and joint beam-atmospheric fits have similar sensitivity  
 3323 to  $\delta_{CP}$  and  $\sin^2(\theta_{23})$  when projected in one-dimension, the benefit of the joint  
 3324 analysis becomes more obvious when the sensitivities are presented in two-

dimensions. The sensitivity of the two fits to the appearance parameters ( $\delta_{CP} - \sin^2(\theta_{13})$ ) are illustrated in Figure 8.35.

The width of the 99% joint fit credible interval in  $\sin^2(\theta_{13})$  is squeezed in the region of  $\delta_{CP} \sim 0$  compared to the beam-only analysis. This is the same behaviour that is seen in the appearance likelihood scans presented in Figure 8.2. The  $1\sigma$  and 90% also exhibit slightly tighter constraints on  $\delta_{CP}$ . This is most prevalent in the region of  $\delta_{CP} \sim 0$  and  $\sin^2(\theta_{13}) \sim 0.03$ . Whilst the atmospheric samples do not have significant sensitivity to  $\sin^2(\theta_{13})$  (as shown in Figure 8.1), they aid in breaking the degeneracy between the oscillation parameters allowing for tighter constraints.

The sensitivity to the disappearance parameters  $\sin^2(\theta_{23}) - \Delta m_{32}^2$ , marginalised over both hierarchies, is presented in Figure 8.36 for both the beam-only and joint beam-atmospheric fits. Whilst the one-dimensional sensitivity comparisons considered so far show the improvements of the joint fit, the two-dimensional projection really shows the benefit of adding the atmospheric samples to the beam samples. The area contained within the IH credible intervals is drastically reduced in the joint fit. This follows from the better determination of the mass hierarchy seen in the Bayes factor comparisons. The  $1\sigma$  joint fit credible interval in the IH region more strongly favours the UO as compared to the beam-only fit. Even in the NH region, the width of the credible intervals in  $\sin^2(\theta_{23})$  decrease, albeit to a smaller extent.

The change in sensitivity to  $\delta_{CP} - \Delta m_{32}^2$  is illustrated in Figure 8.37. As expected, the contours presented within the IH region are much smaller in the joint fit due to the increased sensitivity to mass hierarchy determination. This culminates in a region around  $\delta_{CP} \sim \pi/2$  which is excluded at  $3\sigma$ . This behaviour is not present within the beam-only analysis. Consistent with the previous observations, the area contained within the IH credible intervals is significantly reduced in comparison to the beam-only analysis.

The sensitivity to  $\Delta m_{32}^2$  and  $\sin^2(\theta_{23})$ , as a function of  $\sin^2(\theta_{13})$ , is presented in Figure 8.38 and Figure 8.39, respectively. These sensitivities are marginalised

### Without Reactor Constraint, Both Hierarchies



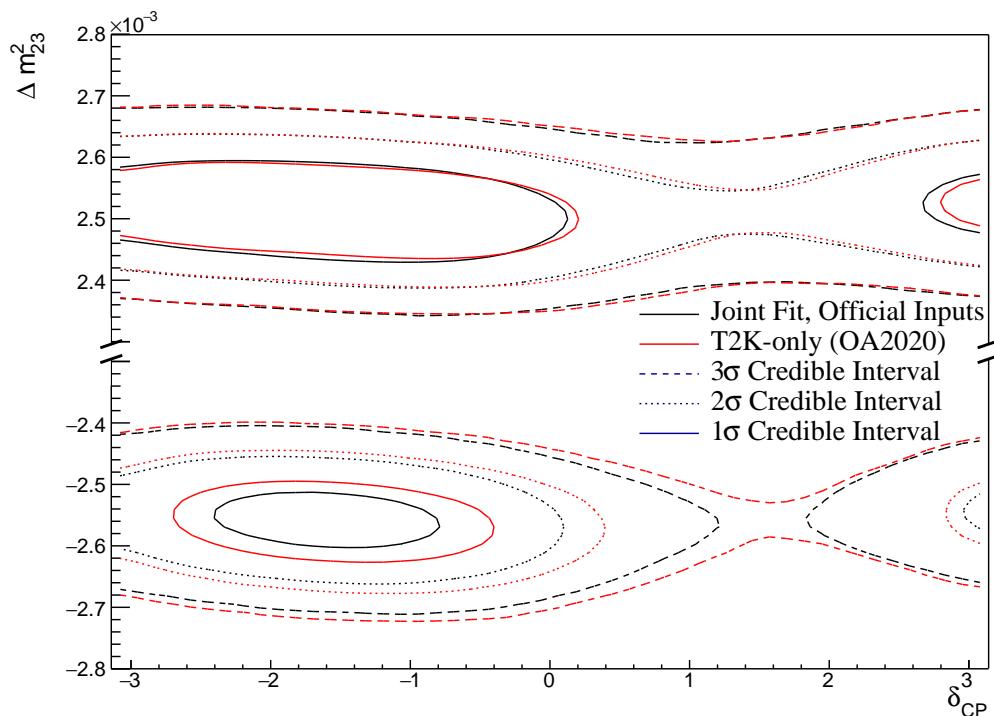
**Figure 8.35:** The two-dimensional posterior probability density distribution in  $\delta_{CP}$  –  $\sin^2(\theta_{13})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

over both hierarchies. As expected from the previous observations, the  $\Delta m_{32}^2$  contours within IH region of the joint fit are much smaller than the beam-only analysis. Notably, the joint fit IH  $1\sigma$  credible intervals exclude the region around the reactor constraint. This is not a bias from the fit as the known value for  $\Delta m_{32}^2$  is in the NH region. This does suggest that the application of the reactor constraint would further increase the preference for NH in the joint fit as compared to its effect on the beam-only analysis.

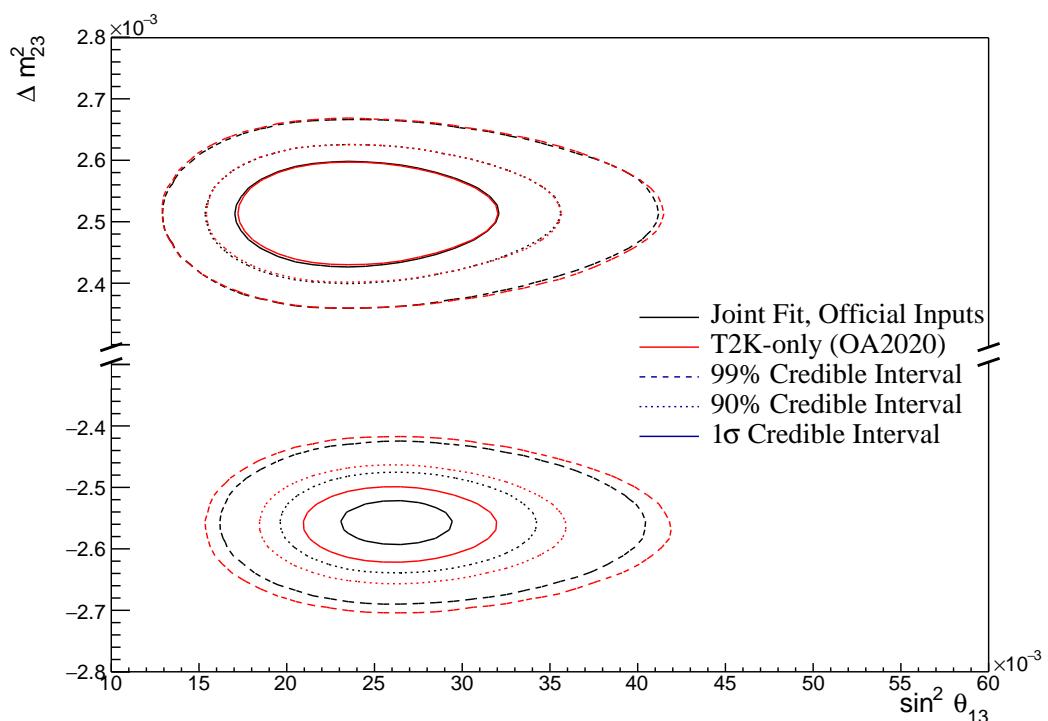
The beam-only and joint beam-atmospheric fits have a slightly different contour shape between the  $\sin^2(\theta_{13})$  and  $\sin^2(\theta_{23})$  parameters, as illustrated by Figure 8.39. The joint analysis disfavours the wrong octant hypothesis more strongly in the region of high  $\sin^2(\theta_{13})$ . This suggests that the application of the reactor constraint will favour the UO more strongly in the joint analysis compared to the beam-only analysis.



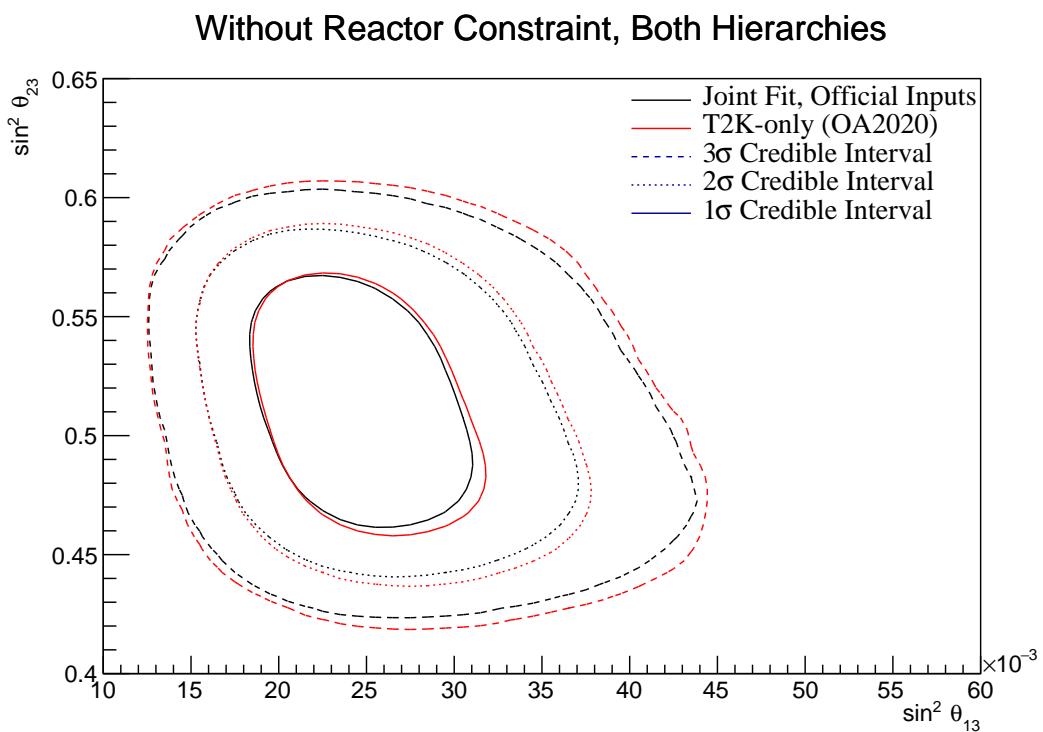
**Figure 8.36:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  –  $\sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.



**Figure 8.37:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \Delta_{CP}$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.



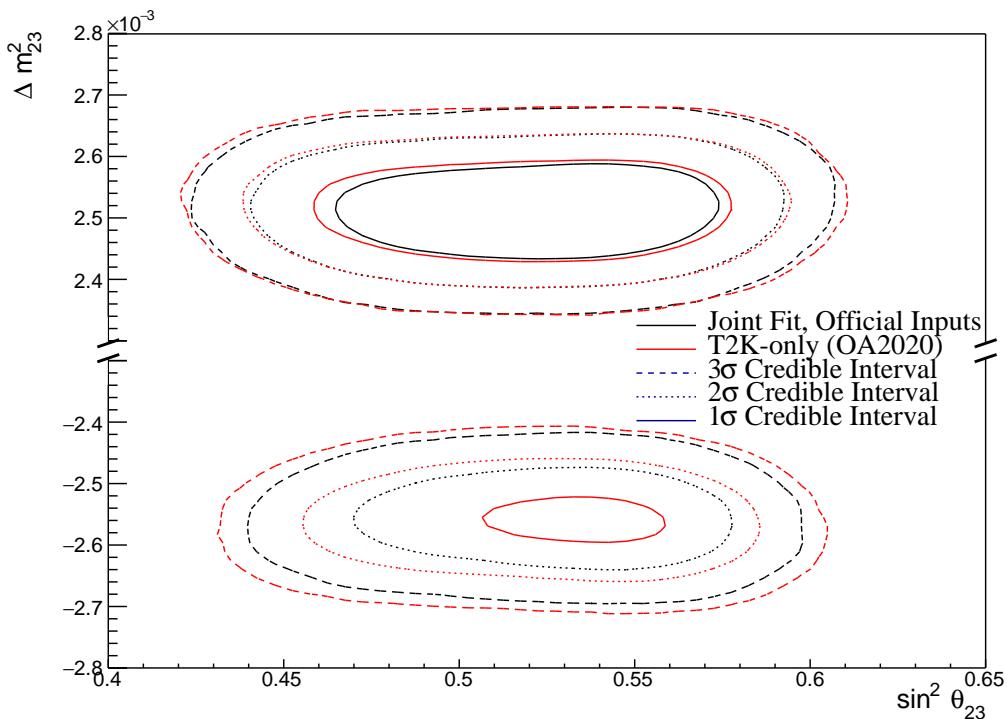
**Figure 8.38:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  –  $\sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.



**Figure 8.39:** The two-dimensional posterior probability density distribution in  $\sin^2(\theta_{23}) - \sin^2(\theta_{13})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

### 3368 8.3.7 Comparison to Latest T2K Sensitivities with Reactor Con- 3369 straint

3370 The comparison between the beam-only and joint beam-atmospheric fits are  
 3371 compared in subsection 8.3.6. Those comparisons were made with the reactor  
 3372 constraint not applied to either of the fits. This section illustrates the com-  
 3373 parison when the reactor constraint is applied. As shown in Figure 8.38, the  
 3374 application of the reactor constraint is expected to significantly increase the  
 3375 joint fit's preference for the NH hypothesis, as compared to the beam-only  
 3376 analysis. Figure 8.40 illustrates the sensitivities of the two fits to the disappearance  
 3377 parameters ( $\sin^2(\theta_{23}) - \Delta m_{32}^2$ ) marginalised over both hierarchies and with the  
 3378 reactor constraint applied. This plot clearly illustrates the benefit of the joint  
 3379 beam and atmospheric analysis. The  $1\sigma$  credible interval in the IH region is  
 3380 entirely removed in the joint analysis, illustrating the improved NH preference.



**Figure 8.40:** The two-dimensional posterior probability density distribution in  $\Delta m_{32}^2 - \sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is applied in both fits. The distributions are marginalised over both hierarchies.

3381 The credible intervals of the joint fit are also tighter in the  $\sin^2(\theta_{23})$  dimension  
3382 than the beam-only analysis in both mass hierarchy regions. This shows that  
3383 beyond the ability of the joint fit to prefer the NH more strongly than the beam-  
3384 only analysis, the precision to which it can measure  $\sin^2(\theta_{23})$  is also improved.  
3385 The Bayes factor for NH preference is calculated as  $B(\text{NH}/\text{IH}) = 7.29$  and  
3386  $B(\text{NH}/\text{IH}) = 3.41$  for the joint beam-atmospheric and beam-only analysis,  
3387 respectively. Whilst both present a significant preference for the NH hypothesis  
3388 (Table 4.1), the joint fit's preference is much stronger. A similar conclusion can be  
3389 made regarding the Bayes factors for UO preference which are  $B(\text{UO}/\text{LO}) = 2.86$   
3390 and  $B(\text{UO}/\text{LO}) = 2.67$  for the joint beam-atmospheric and beam-only analysis,  
3391 respectively. Both of these represent a mild preference for the UO but there is  
3392 a stronger preference observed in the joint analysis.

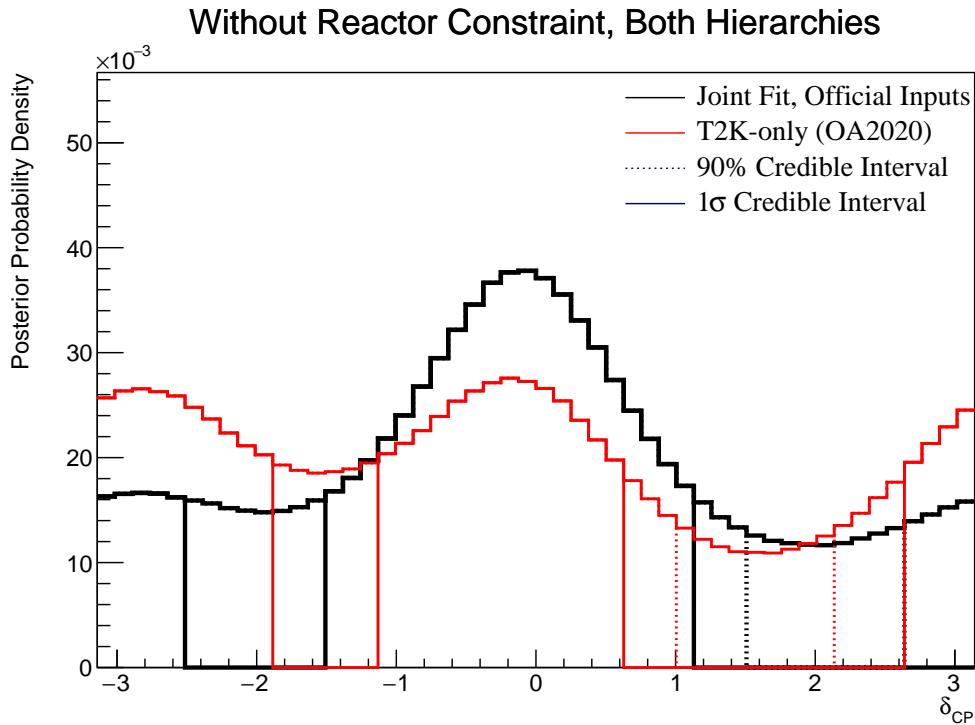
3393 The sensitivity of the beam-only and joint beam-atmospheric analyses, to the  
3394 appearance parameters ( $\delta_{CP} - \sin^2(\theta_{13})$ ), are compared in Figure 8.41. These  
3395 results are marginalised over both hierarchies and include the reactor constraint  
3396 on  $\sin^2(\theta_{13})$ . For this particular set of known oscillation parameters (AsimovA  
3397 defined in Table 2.2), the beam-only analysis dominates the sensitivity. The  
3398 joint fit does slightly increase the sensitivity to  $\delta_{CP}$  but it does not change any  
3399 conclusions that would be made.



**Figure 8.41:** The two-dimensional posterior probability density distribution in  $\delta_{CP}$  –  $\sin^2(\theta_{13})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is applied in both fits. The distributions are marginalised over both hierarchies.

### 3400 8.3.8 Effect of Asimov Parameter Set

3401 Figure 8.4 and Figure 8.5 show that the choice of the parameter set at which the  
 3402 Asimov data is made can affect the conclusion. ‘AsimovA’ oscillation parameters  
 3403 are defined at a region of  $\delta_{CP}$  which is dominated by the T2K experiment. This  
 3404 explains why the addition of the atmospheric samples does not significantly in-  
 3405 crease the sensitivity to  $\delta_{CP}$ , as illustrated in subsection 8.3.6 and subsection 8.3.7.  
 3406 This section presents the sensitivities when ‘AsimovB’ oscillation parameters,  
 3407 as defined in Table 2.2, are assumed (alongside the post-BANFF tune) when  
 3408 building the Asimov data.



**Figure 8.42:** The one-dimensional posterior probability density distribution in  $\delta_{CP}$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3409 The sensitivity to  $\delta_{CP}$  for the joint beam and atmospheric fit is presented  
 3410 in Figure 8.42. The results are compared to those from the beam-only analysis  
 3411 in [182]. The reactor constraint is not applied in either of the fits. The known  
 3412 oscillation parameter value is  $\delta_{CP} = 0$ . The shape of the posterior distribution

3413 from the joint analysis is more peaked at  $\delta_{CP} = 0$  as compared to the beam-only  
3414 analysis which has approximately the same posterior probability density at  $\delta_{CP} =$   
3415 0 and  $\delta_{CP} = \pi$ . This shows the ability of the joint analysis to better determine the  
3416 correct phase of  $\delta_{CP}$  if the true value was CP-conserving. The  $1\sigma$  credible intervals  
3417 and the position of the highest posterior probability density are given in Table 8.9.

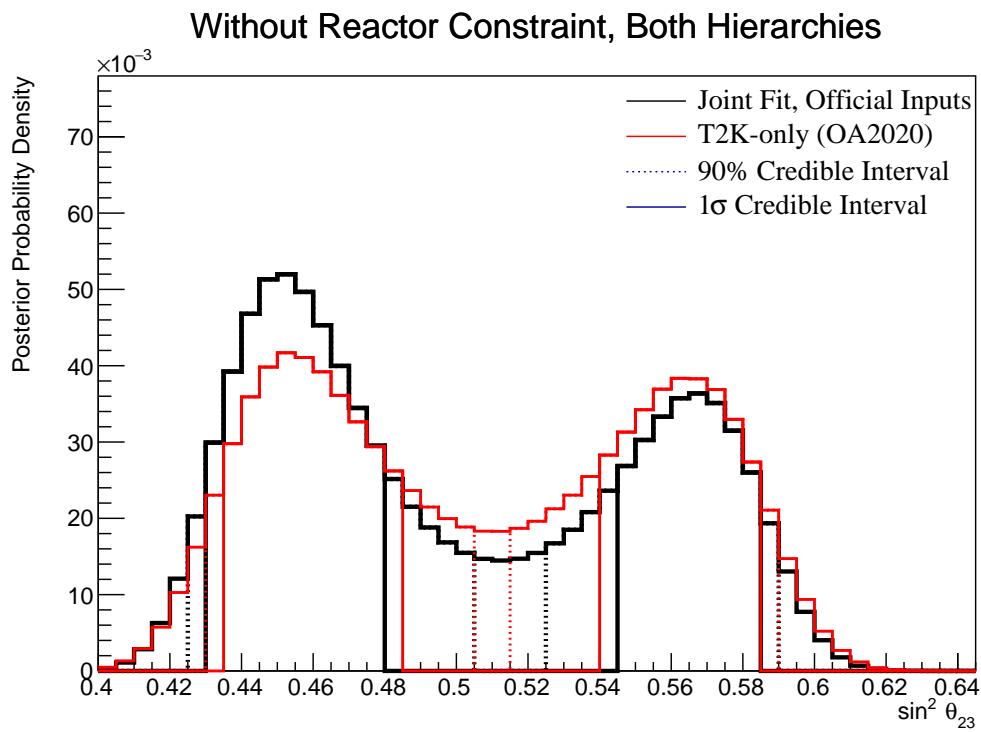
Parameter	Interval	HPD
$\delta_{CP}$ , (BH)	$[-\pi, -2.51], [-1.51, 1.31]$	-0.06
$\delta_{CP}$ , (NH)	$[-1.13, 1.63]$	0.06
$\delta_{CP}$ , (IH)	$[-3.02, -1.88], [-1.76, 0.13]$	-0.44
$\Delta m_{32}^2$ (BH) [ $\times 10^{-3}\text{eV}^2$ ]	$[-2.60, -2.49], [2.46, 2.59]$	2.51
$\Delta m_{32}^2$ (NH) [ $\times 10^{-3}\text{eV}^2$ ]	$[2.47, 2.56]$	2.52
$\Delta m_{32}^2$ (IH) [ $\times 10^{-3}\text{eV}^2$ ]	$[-2.61, -2.52]$	-2.57
$\sin^2(\theta_{23})$ (BH)	$[0.43, 0.48], [0.55, 0.59]$	0.45
$\sin^2(\theta_{23})$ (NH)	$[0.43, 0.49], [0.55, 0.58]$	0.45
$\sin^2(\theta_{23})$ (IH)	$[0.44, 0.48], [0.54, 0.59]$	0.57

**Table 8.9:** The position of the highest posterior probability density (HPD) and width of the  $1\sigma$  credible interval for the SK atmospheric-only fit. The reactor constraint is not applied. The values are presented by which hierarchy hypothesis is assumed: marginalised over both hierarchies (BH), normal hierarchy only (NH) and inverted hierarchy only (IH).

3418 Naively, if just the  $1\sigma$  credible interval were considered without observing  
3419 the shape of the distribution, it would appear that the joint analysis would  
3420 have a worse sensitivity to  $\delta_{CP}$  due to the larger interval around  $\delta_{CP}$ . The  
3421  $1\sigma$  credible interval for the beam-only analysis is given as the range  $\delta_{CP} =$   
3422  $[-\pi, -1.88], [-1.13, 0.63]$  and  $[2.64, \pi]$  which contains 56% of all values of  $\delta_{CP}$ .  
3423 The joint beam and atmospheric analysis contains 52% of all  $\delta_{CP}$  values within  
3424 the  $1\sigma$  credible interval. Therefore, if the area within the  $1\sigma$  credible interval  
3425 were to be compared between the two fits, the joint analysis would be shown  
3426 to have better precision.

3427 This contradiction stems from the methodology in which the credible interval  
3428 is calculated. The technique used in this analysis (documented in subsection 4.3.2)

3429 fills the credible interval by selecting bins in order of magnitude until 68% of the  
3430 posterior density is contained. If instead, the credible interval was calculated  
3431 by expanding around the highest posterior probability, the benefits of the joint  
3432 fit would be more obvious. In the case where the shape of the posterior was  
3433 Gaussian, these two techniques would be equivalent.



**Figure 8.43:** The one-dimensional posterior probability density distribution in  $\sin^2(\theta_{23})$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

3434 The sensitivity of the joint beam and atmospheric fit to  $\sin^2(\theta_{23})$  is presented  
3435 in Figure 8.43. The sensitivity is compared to that of the beam-only analysis  
3436 in [182]. The reactor constraint is not applied in either of the fits being com-  
3437 pared. The Asimov parameter value is  $\sin^2(\theta_{23}) = 0.45$  and the sensitivities are  
3438 marginalised over both hierarchies. Clearly, the joint beam and atmospheric  
3439 fit has a much larger probability density in the region surrounding the known  
3440 oscillation parameters. This shows the better octant determination of the joint  
3441 analysis compared to the beam-only fit. The ratio of the posterior density at

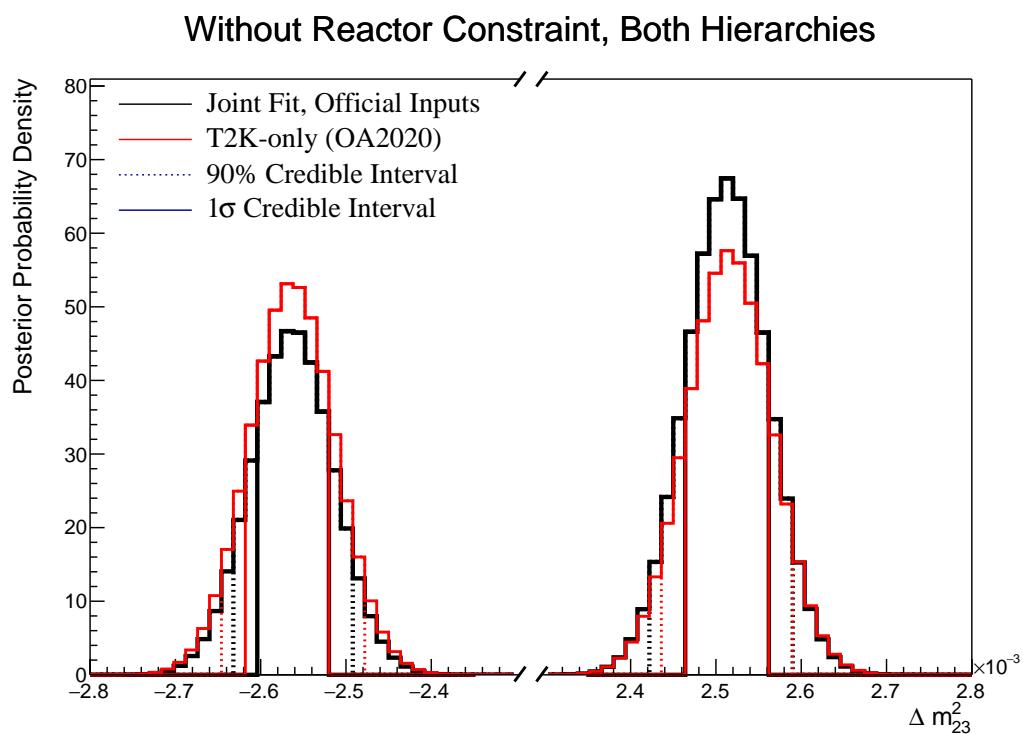
3442 the peak of the lower octant to the peak of the upper octant from the joint fit is  
3443 1.43 compared to 1.09 from the beam-only analysis. This shows further support  
3444 for the joint analysis in correctly selecting the lower octant, which is the correct  
3445 hypothesis given the known oscillation parameters.

	LO ( $\sin^2 \theta_{23} < 0.5$ )	UO ( $\sin^2 \theta_{23} > 0.5$ )	Sum
NH ( $\Delta m_{32}^2 > 0$ )	0.35	0.24	0.59
IH ( $\Delta m_{32}^2 < 0$ )	0.19	0.22	0.41
Sum	0.54	0.46	1.00

**Table 8.10:** The distribution of steps in a joint beam and atmospheric fit, presented as the fraction of steps in the upper (UO) and lower (LO) octants and the normal (NH) and inverted (IH) hierarchies. The reactor constraint is not applied. The Bayes factors are calculated as  $B(\text{NH}/\text{IH}) = 1.43$  and  $B(\text{LO}/\text{UO}) = 1.19$ .

3446 The distribution of steps, split by hierarchy and octant hypothesis, is presented  
3447 in Table 8.10. The Bayes factor for hierarchy and octant determination are  
3448  $B(\text{NH}/\text{IH}) = 1.43$  and  $B(\text{LO}/\text{UO}) = 1.19$ , respectively. The octant Bayes factor  
3449 is now presented as LO/UO as the known oscillation parameter is contained  
3450 within the lower octant. These values compare to  $B(\text{NH}/\text{IH}) = 1.08$  and  
3451  $B(\text{LO}/\text{UO}) = 0.91$  from the beam-only analysis. This shows additional evidence  
3452 of the joint analysis's preference for selecting the correct octant and hierarchy  
3453 hypothesis. Comparisons to the AsimovA Bayes factors presented in Table 8.6  
3454 show how the preference for the correct octant and hierarchy depend on the  
3455 true value of  $\delta_{CP}$  and  $\sin^2(\theta_{23})$ .

3456 The sensitivity of the beam-only and joint beam-atmospheric analysis to  
3457  $\Delta m_{32}^2$  is given in Figure 8.44. Both of the results are marginalised over both  
3458 hierarchies and the reactor constraint is not applied in either analysis. The joint  
3459 analysis has a stronger preference for the correct hierarchy (NH) which is shown  
3460 by the higher Bayes factor ( $B(\text{NH}/\text{IH}) = 1.43$ ) compared to the beam-only  
3461 analysis ( $B(\text{NH}/\text{IH}) = 1.08$ ).



**Figure 8.44:** The one-dimensional posterior probability density distribution in  $\Delta m_{32}^2$  compared between the joint beam-atmospheric fit (Black) and the latest T2K sensitivities (Red) [182]. The reactor constraint is not applied in either fit. The distributions are marginalised over both hierarchies.

# 9

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## Conclusions and Outlook

# Appendices

# A

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3466

## Atmospheric Sample Spectra

3467 This appendix documents the interaction mode breakdown of all the atmospheric  
3468 samples used within the analysis. The generated tune of the model parameters  
3469 and the Asimov A oscillation parameter set (defined in Table 2.2) are assumed.  
3470 The livetime of SK-IV is taken to be 3244.4 days.

### 3471 A.1 Binning

3472 The lepton momentum and cosine zenith binning edges for the atmospheric  
3473 samples used within this analysis are defined in Table A.1.

### 3474 A.2 Fully Contained Sub-GeV Samples

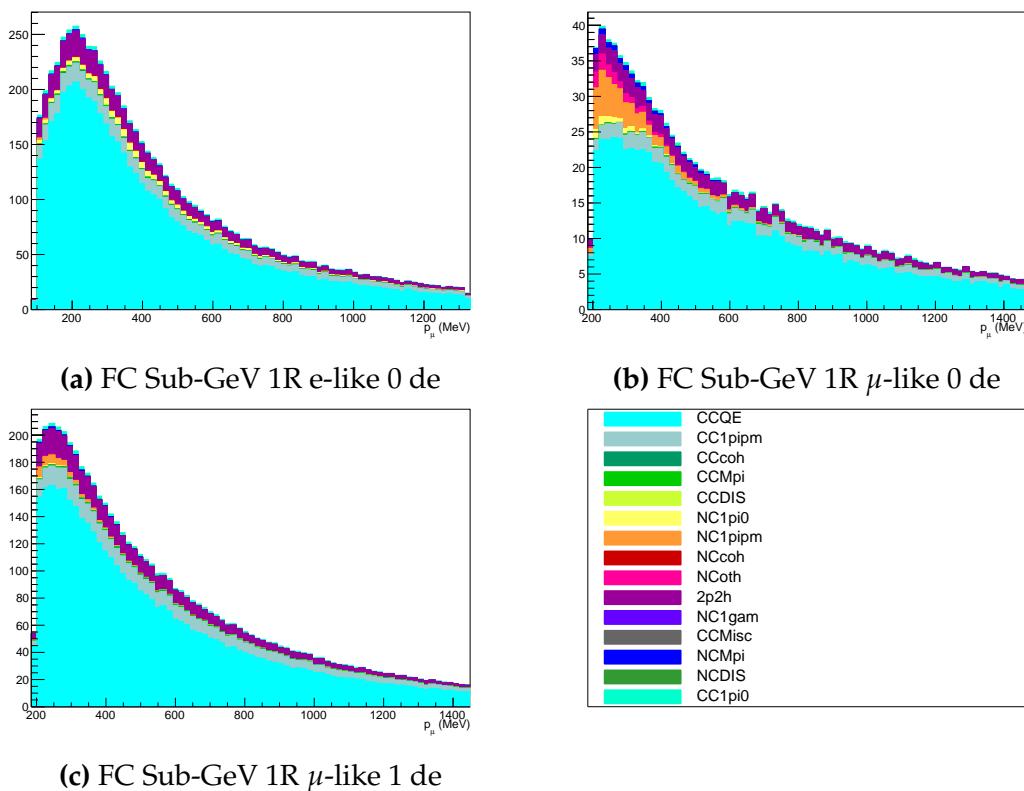
3475 The interaction mode breakdown of the fully contained Sub-GeV samples are  
3476 shown in Figure A.1 and Figure A.2, for the samples with enriched CC0 $\pi$  and  
3477 CC1 $\pi^\pm$  respectively.

3478 The CC0 $\pi$  sample are dominated by CCQE events ( $\sim 70\%$ ) with smaller  
3479 contributions of 2p2h ( $\sim 12\%$ ) and CC1 $\pi$  ( $\sim 10\%$ ) components. The energy peaks  
3480 around 300 MeV, which is slightly below that of the T2K samples but still has  
3481 significant contribution upto 1 GeV which overlaps the T2K sample energy range.

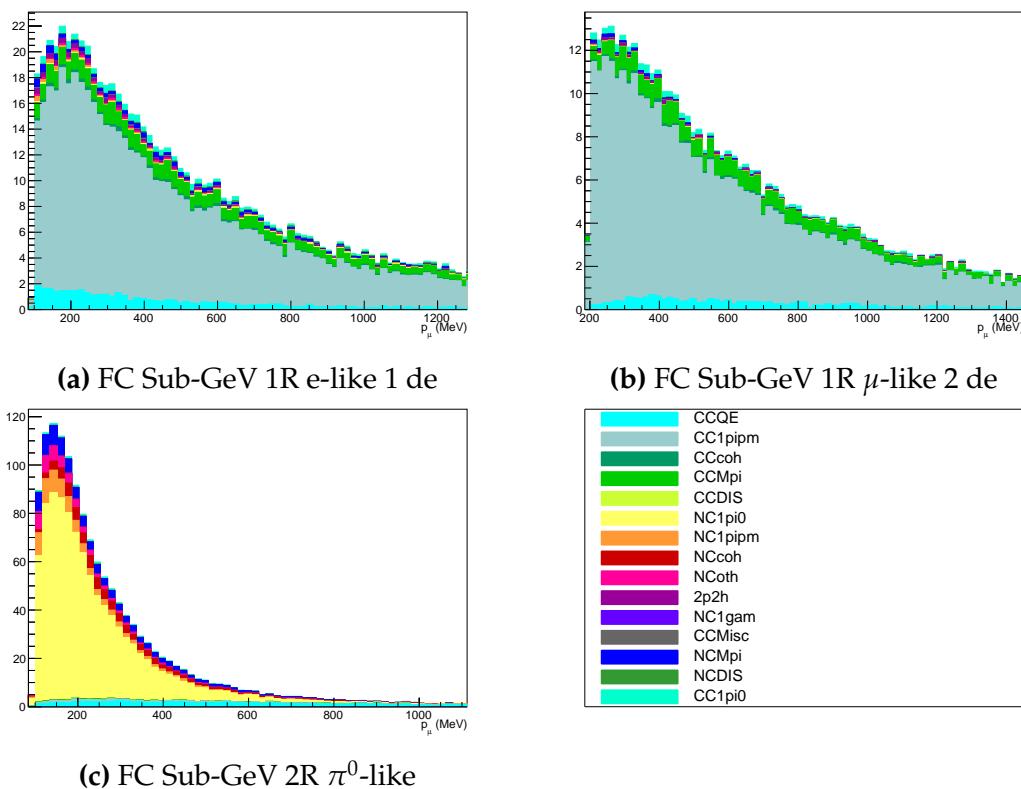
Sample	$\cos(\theta_Z)$ Bins	Momentum Bin Edges ( $\log_{10}(P)$ MeV)
SubGeV-elike-0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-elike-1dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-1dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-mulike-2dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV-pi0like	1	2.0, 2.2, 2.4, 2.6, 2.8, 3.2
MultiGeV-elike-nue	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV-elike-nuebar	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV-mulike	10	3.0, 3.4, 5.0
MultiRing-elike-nue	10	3.0, 3.4, 3.7, 5.0
MultiRing-elike-nuebar	10	3.0, 3.4, 3.7, 5.0
MultiRing-mulike	10	2.0, 3.124, 3.4, 3.7, 5.0
MultiRing-Other1	10	3.0, 3.4, 3.7, 4.0, 5.0
PC-Stop	10	2.0, 3.4, 5.0
PC-Through	10	2.0, 3.124, 3.4, 3.7, 5.0
Upmu-Stop	10	3.2, 3.4, 3.7, 8.0
Upmu-Through-Showering	10	2.0, 8.0
Upmu-Through-NonShowering	10	2.0, 8.0

**Table A.1:** The reconstructed cosine zenith and lepton momentum binning assigned to the atmospheric samples. The “ $\cos(\theta_Z)$  Bins” column illustrates the number of bins uniformly distributed over the  $-1.0 \leq \cos(\theta_Z) \leq 1.0$  region for fully and partially contained samples and  $-1.0 \leq \cos(\theta_Z) \leq 0.0$  region for up- $\mu$  samples.

3482      The one-ring CC1 $\pi$  samples, where the pion is tagged via its decay electron,  
 3483      are dominated by CC1 $\pi$  events ( $\sim 75\%$ ) with a small contribution of CCM $\pi$   
 3484      ( $\sim 10\%$ ). The two-ring pion sample is mostly dominated by the NC1 $\pi^0$  via  
 3485      resonances, and has several equally-sized contributions from CCQE, NC1 $\pi^\pm$  via  
 3486      resonances, and NC coherent pion production, where the  $\pi^0$  likely comes from  
 3487      nucleon and  $\pi^\pm$  final state interactions in the nucleus.



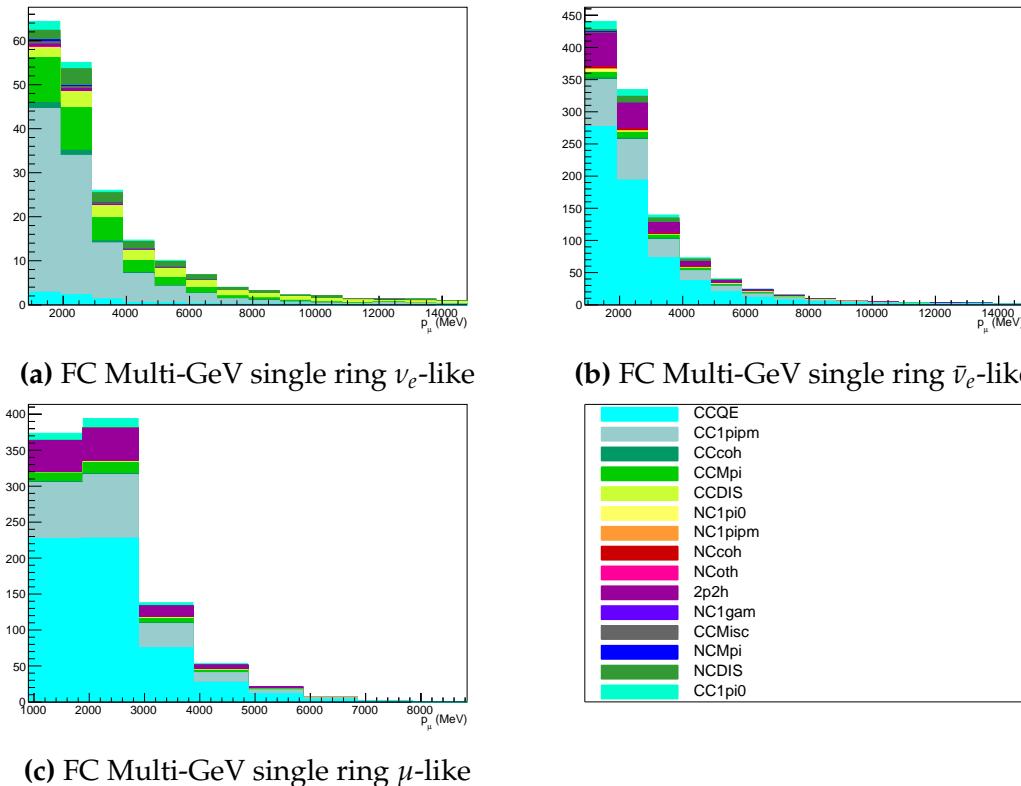
**Figure A.1:** Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting CC $0\pi$  events.



**Figure A.2:** Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting single pion events.

### 3488 A.3 Fully Contained Multi-GeV Samples

3489 The interaction mode breakdown of fully contained multi-GeV samples is high-  
 3490 lighted in Figure A.3. Due to the event selection applied in SK which targets  $\pi^+$   
 3491 and  $\pi^-$  separation, the  $\nu_e$  sample mainly consists of events with pions (single pion  
 3492 production or multi-pion/DIS interactions). The pion separation is explained in  
 3493 Section section 6.1. This reasoning also explains the significant CCQE contribution  
 3494 of the  $\bar{\nu}_e$  sample. The muon-like sample is dominated by CCQE interactions with  
 3495  $\sim 10 - 15\%$  2p2h and CC1 $\pi$  contribution of events.



**Figure A.3:** Breakdown by interaction mode of the FC Multi-GeV single ring atmospheric samples.

## 3496 A.4 Fully Contained Multi-Ring Samples

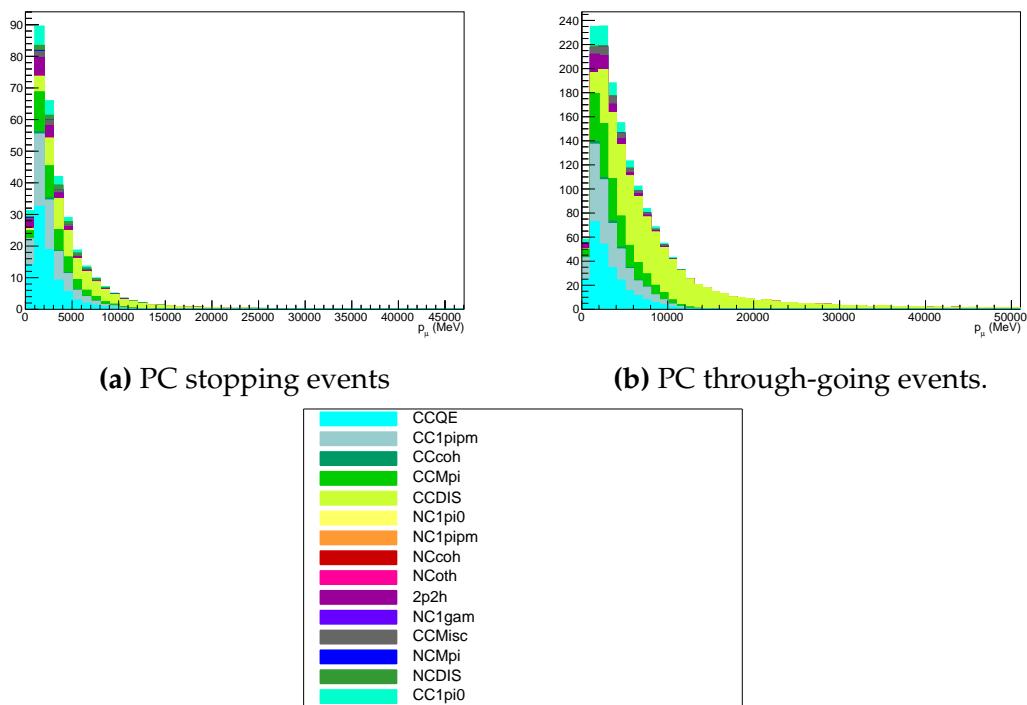
3497 The interaction mode breakdown of fully contained multi-ring events is shown  
 3498 in Figure A.4. These samples see more interaction modes contributing in general,  
 3499 and there is a much larger contribution from multi-pion and DIS interaction  
 3500 modes, compared to the other samples.



**Figure A.4:** Breakdown by interaction mode of the FC Multi-GeV multi-ring atmospheric samples.

### 3501 A.5 Partially Contained Samples

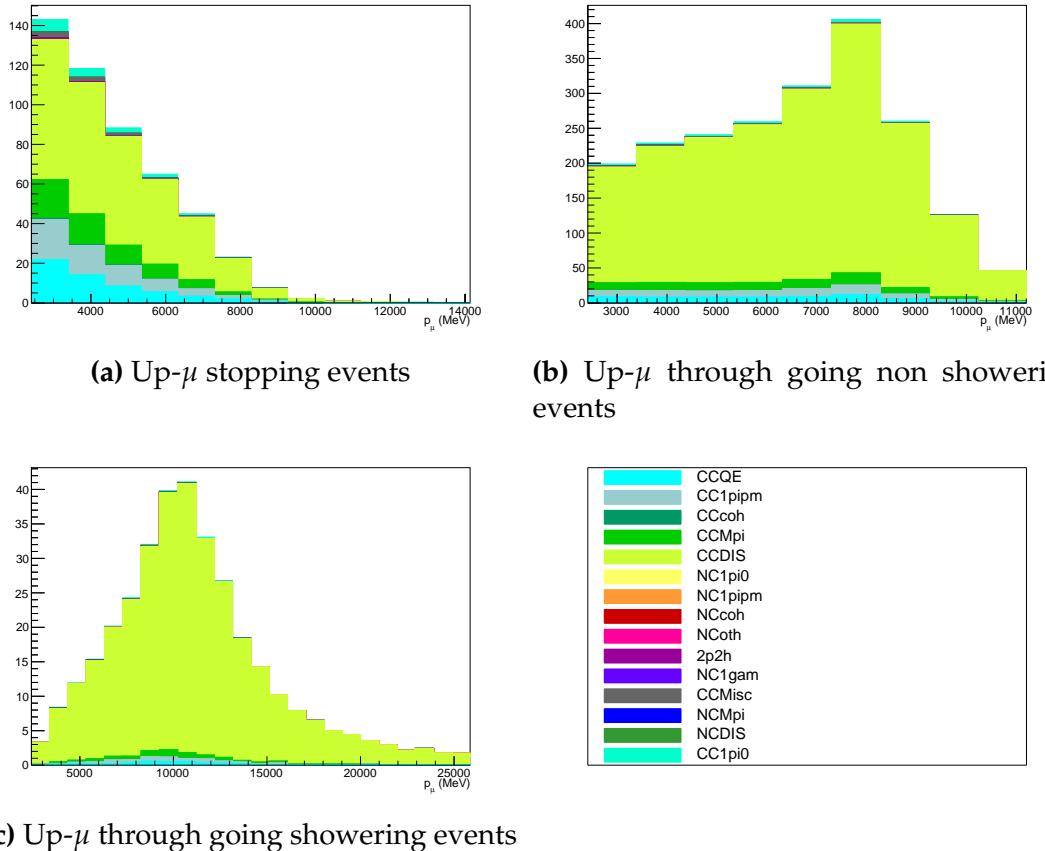
3502 The breakdown for partially contained samples is highlighted in Figure A.5.  
 3503 As with the multi-ring samples, there is no dominating interaction mode. The  
 3504 neutrino energies of events in this sample extend into the tens of GeV and become  
 3505 dominated by DIS interaction modes in the high energy limit.



**Figure A.5:** Breakdown by interaction mode of the PC atmospheric samples.

## 3506 A.6 Upward-Going Muon Samples

3507 The breakdown for upward-going muons is illustrated in Figure A.6. These  
 3508 samples are significantly dominated by DIS interactions with energies extending  
 3509 up into the hundreds of GeV.



**Figure A.6:** Breakdown by interaction mode of the atmospheric upward going muon samples.

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