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# Simultaneous Measurement of Oscillation Parameters in Beam and Atmospheric Neutrino Data from Tokai-to-Kamioka and Super-Kamiokande Experiments

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Version 1.2

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A Dissertation Submitted to Oxford University  
11 for the Degree of Doctor of Philosophy

13                   **Simultaneous Measurement of**

14                   **Oscillation Parameters in Beam and**

15                   **Atmospheric Neutrino Data from**

16                   **Tokai-to-Kamioka and**

17                   **Super-Kamiokande Experiments**

18                   *Abstract*

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## Acknowledgements

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<sup>147</sup> **Chapter 1**

<sup>148</sup> **Introduction**

<sup>149</sup> **Chapter 2**

<sup>150</sup> **Neutrino Oscillation Physics**

<sup>151</sup> When first proposed, neutrinos were expected to be massless fermions that only in-  
<sup>152</sup> teract through weak and gravitational forces. This meant they were very difficult to  
<sup>153</sup> detect as they can pass through significant amounts of matter without interacting. De-  
<sup>154</sup> spite this, experimental neutrino physics has developed with many different detection  
<sup>155</sup> techniques and neutrino sources being used today. In direct tension with standard  
<sup>156</sup> model physics, neutrinos have been determined to oscillate between different lepton  
<sup>157</sup> flavours, requiring them to have mass.

<sup>158</sup> The observation techniques which lead to the discovery of the neutrino are doc-  
<sup>159</sup> umented in section 2.1. The theory underpinning neutrino oscillation is described  
<sup>160</sup> in section 2.2 and includes the approximations which can be made to simplify the  
<sup>161</sup> understanding of neutrino oscillation in the two-flavour approximation. Past, current,  
<sup>162</sup> and future neutrino experiments are detailed in section 2.3, including the reactor,  
<sup>163</sup> atmospheric, and long-baseline accelerator neutrino sources that have been used to  
<sup>164</sup> successfully constrain oscillation parameters. Finally, the current state of oscillation  
<sup>165</sup> parameter measurements are summarised in section 2.4.

<sup>166</sup> **2.1. Discovery of Neutrinos**

<sup>167</sup> At the start of the 20<sup>th</sup> century, the electrons emitted from the  $\beta$ -decay of the nucleus  
<sup>168</sup> were found to have a continuous energy spectrum [1,2]. This observation seemingly  
<sup>169</sup> broke the energy conservation invoked within that period's nuclear models. Postulated

170 in 1930 by Pauli as the solution to this problem, the neutrino (originally termed  
171 “neutron”) was theorized to be an electrically neutral spin-1/2 fermion with a mass of  
172 the same order of magnitude as the electron [3]. This neutrino was to be emitted with  
173 the electron in  $\beta$ -decay to alleviate the apparent breaking of energy conservation. As a  
174 predecessor of today’s weak interaction model, Fermi’s theory of  $\beta$ -decay developed  
175 the understanding by coupling the four constituent particles; electron, proton, neutron,  
176 and neutrino, into a consistent model [4].

177 Whilst Pauli was not convinced of the ability to detect neutrinos, the first observa-  
178 tions of the particle were made in the mid-1950s when neutrinos from a reactor were  
179 observed via the inverse  $\beta$ -decay (IBD) process,  $\bar{\nu}_e + p \rightarrow n + e^+$  [5, 6]. The detector  
180 consisted of two parts: a neutrino interaction medium and a liquid scintillator. The  
181 interaction medium was built from two water tanks. These were loaded with cadmium  
182 chloride to allow increased efficiency of neutron capture. The positron emitted from  
183 IBD annihilates,  $e^+ + e^- \rightarrow 2\gamma$ , generating a prompt signal and the neutron is captured  
184 on the cadmium via  $n + {}^{108}Cd \rightarrow {}^{109*}Cd \rightarrow {}^{109}Cd + \gamma$ , producing a delayed signal. An  
185 increase in the coincidence rate was observed when the reactor was operating which  
186 was interpreted as interactions from neutrinos generated in the reactor.

187 After the discovery of the  $\nu_e$ , the natural question of how many flavours of neutrino  
188 exist was asked. In 1962, a measurement of the  $\nu_\mu$  was conducted at the Brookhaven  
189 National Laboratory [7]. A proton beam was directed at a beryllium target, generating  
190 a  $\pi$ -dominated beam which then decayed via  $\pi^\pm \rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu)$ , and the subsequent  
191 interactions of the  $\nu_\mu$  were observed. As the subsequent interaction of the neutrino  
192 generates muons rather than electrons, it was determined the  $\nu_\mu$  was fundamentally  
193 different from  $\nu_e$ . The final observation to be made was that of the  $\nu_\tau$  from the DONUT  
194 experiment [8]. Three neutrinos seem the obvious solution as it mirrors the known  
195 number of charged lepton (as they form weak isospin doublets) but there could be

196 evidence of more. Several neutrino experiments have found anomalous results [9, 10]  
197 which could be attributed to sterile neutrinos. However, cosmological observations  
198 indicate the number of neutrino species  $N_{eff} = 2.99 \pm 0.17$  [11], as measured from  
199 the cosmic microwave background power spectrum, and LEP measured the number  
200 of active neutrino flavours to be  $N_\nu 2.9840 \pm 0.0082$  [12] from measurements of the  
201  $Z$ -decay width.

## 202 2.2. Theory of Neutrino Oscillation

203 As direct evidence of beyond Standard Model physics, a neutrino generated with  
204 lepton flavour  $\alpha$  can change into a different lepton flavour  $\beta$  after propagating some  
205 distance. This phenomenon is called neutrino oscillation and requires that neutrinos  
206 must have a non-zero mass (as seen in subsection 2.2.1). This observation is direct  
207 evidence of beyond standard model physics. This behaviour has been characterised  
208 by the Pontecorvo-Maki-Nakagawa-Sakata (PMNS) [13–15] mixing matrix which  
209 describes how the flavour and mass of neutrinos are associated. This is analogous to  
210 the Cabibbo-Kobayashi-Maskawa (CKM) [16] matrix measured in quark physics.

### 211 2.2.1. Three Flavour Oscillations

212 The PMNS parameterisation defines three flavour eigenstates,  $\nu_e$ ,  $\nu_\mu$  and  $\nu_\tau$  (indexed  
213  $\nu_\alpha$ ), which are eigenstates of the weak interaction and three mass eigenstates,  $\nu_1$ ,  $\nu_2$  and  
214  $\nu_3$  (indexed  $\nu_i$ ). Each mass eigenstate is the superposition of all three flavour states,

$$|\nu_i\rangle = \sum_\alpha U_{\alpha i} |\nu_\alpha\rangle. \quad (2.1)$$

215 Where  $U$  is the PMNS matrix which is unitary and connects the mass and flavour

216 eigenstates.

217 The weak interaction couples to flavour eigenstates so neutrinos interact with  
218 leptons of the same flavour. The propagation of a neutrino flavour eigenstate, in a  
219 vacuum, can be re-written as a plane-wave solution to the time-dependent Schrödinger  
220 equation,

$$|\nu_\alpha(t)\rangle = \sum_i U_{\alpha i}^* |\nu_i\rangle e^{-i\phi_i}. \quad (2.2)$$

221 The probability of observing a neutrino of flavour eigenstate  $\beta$  from one which  
222 originated as flavour  $\alpha$  can be calculated as,

$$P(\nu_\alpha \rightarrow \nu_\beta) = |\langle \nu_\beta | \nu_\alpha(t) \rangle|^2 = \sum_{i,j} U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* e^{-i(\phi_j - \phi_i)} \quad (2.3)$$

223 The  $\phi_i$  term can be expressed in terms of the energy,  $E_i$ , and magnitude of the  
224 three momenta,  $p_i$ , of the neutrino,  $\phi_i = E_i t - p_i x$  ( $t$  and  $x$  being time and position  
225 coordinates). Therefore,

$$\phi_j - \phi_i = E_j t - E_i t - p_j x + p_i x. \quad (2.4)$$

226 For a relativistic particle,  $E_i \gg m_i$ ,

$$p_i = \sqrt{E_i^2 - m_i^2} \approx E_i - \frac{m_i^2}{2E_i}. \quad (2.5)$$

<sup>227</sup> Making the approximations that neutrinos are relativistic, the mass eigenstates  
<sup>228</sup> were created with the same energy and that  $x = L$ , where  $L$  is the distance traveled by  
<sup>229</sup> the neutrino, Equation 2.4 then becomes

$$\phi_j - \phi_i = \frac{\Delta m_{ij}^2 L}{2E}, \quad (2.6)$$

<sup>230</sup> where  $\Delta m_{ij}^2 = m_j^2 - m_i^2$ . This, combined with further use of unitarity relations  
<sup>231</sup> results in Equation 2.3 becoming

$$P(\nu_\alpha \rightarrow \nu_\beta) = \delta_{\alpha\beta} - 4 \sum_{i>j} \Re \left( U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin^2 \left( \frac{\Delta m_{ij}^2 L}{4E} \right) + (-) 2 \sum_{i>j} \Im \left( U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^* \right) \sin \left( \frac{\Delta m_{ij}^2 L}{2E} \right). \quad (2.7)$$

<sup>232</sup> Where  $\delta_{\alpha\beta}$  is the Kronecker delta function and the negative sign on the last term is  
<sup>233</sup> included for the oscillation probability of antineutrinos.

<sup>234</sup> Typically, the PMNS matrix is parameterised into three mixing angles, a charge  
<sup>235</sup> parity (CP) violating phase  $\delta_{CP}$ , and two Majorana phases  $\alpha_{1,2}$ ,

$$U = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\text{Atmospheric, Accelerator}} \underbrace{\begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta_{CP}} \\ 0 & 1 & 0 \\ -s_{13}e^{-i\delta_{CP}} & 0 & c_{13} \end{pmatrix}}_{\text{Reactor, Accelerator}} \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Reactor, Solar}} \underbrace{\begin{pmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{Majorana}}. \quad (2.8)$$

<sup>236</sup> Where  $s_{ij} = \sin(\theta_{ij})$  and  $c_{ij} = \cos(\theta_{ij})$ . The oscillation parameters are often  
<sup>237</sup> grouped; (1,2) as “solar”, (2,3) as “atmospheric” and (1,3) as “reactor”. Many  
<sup>238</sup> neutrino experiments aim to measure the PMNS parameters from a wide array of  
<sup>239</sup> origins, as is the purpose of this thesis.

<sup>240</sup> The Majorana phase,  $\alpha_{1,2}$ , included within the fourth matrix in Equation 2.8 is only  
<sup>241</sup> included for completeness. For an oscillation analysis experiment, any terms contain-  
<sup>242</sup> ing this phase disappear due to taking the expectation value of the PMNS matrix.  
<sup>243</sup> Measurements of these phases are typically performed by experiments searching for  
<sup>244</sup> neutrino-less double  $\beta$ -decay [17].

<sup>245</sup> A two flavour approximation can be obtained when one assumes the third mass  
<sup>246</sup> eigenstate is degenerate with another. As discussed in section 2.3, it is found that  
<sup>247</sup>  $\Delta m_{21}^2 \ll |\Delta m_{31}^2|$ . This results in the two flavour approximation being reasonable for  
<sup>248</sup> understanding the features of the oscillation. In this two flavour case, the mixing  
<sup>249</sup> matrix becomes,

$$U_{2 \text{ Flav.}} = \begin{pmatrix} \cos(\theta) & \sin(\theta) \\ -\sin(\theta) & \cos(\theta) \end{pmatrix}. \quad (2.9)$$

250 This culminates in the oscillation probability,

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\alpha) &= 1 - \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right), \\ P(\nu_\alpha \rightarrow \nu_\beta) &= \sin^2(2\theta) \sin^2\left(\frac{\Delta m^2 L}{4E}\right). \end{aligned} \quad (2.10)$$

251 Where  $\alpha \neq \beta$ . For a fixed neutrino energy, the oscillation probability is a sinusoidal  
252 function depending upon the distance over which the neutrino propagates. The  
253 frequency and amplitude of oscillation are dependent upon  $\Delta m^2/4E$  and  $\sin^2 2\theta$ ,  
254 respectively. The oscillation probabilities presented thus far assume  $c = 1$ , where  
255  $c$  is the speed of light in vacuum. In more familiar units, the maximum oscillation  
256 probability for a fixed value of  $\theta$  is given at  $L[km]/E[GeV] \sim 1.27/\Delta m^2$ . It is this  
257 calculation that determines the best  $L/E$  value for a given experiment to be designed  
258 around for measurements of a specific value of  $\Delta m^2$ .

### 259 2.2.2. The MSW Effect

260 The theory of neutrino oscillation in a vacuum has been described in subsection 2.2.1.  
261 However, the beam neutrinos and atmospheric neutrinos originating from below the  
262 horizon propagate through matter in the Earth. The coherent scattering of neutrinos  
263 from a material target modifies the Hamiltonian of the system. This results in a change  
264 in the oscillation probability. Notably, charged current scattering ( $\nu_e + e^- \rightarrow \nu_e + e^-$ ,  
265 propagated by a  $W$  boson) only affects electron neutrinos whereas the neutral current  
266 scattering ( $\nu_l + l^- \rightarrow \nu_l + l^-$ , propagated by a  $Z^0$  boson) interacts through all neutrino  
267 flavours equally. In the two-flavour approximation, the effective mixing parameter  
268 becomes

$$\sin^2(2\theta) \rightarrow \sin^2(2\theta_m) = \frac{\sin^2(2\theta)}{(A/\Delta m^2 - \cos(2\theta))^2 + \sin^2(2\theta)}, \quad (2.11)$$

269 where  $A = 2\sqrt{2}G_F N_e E$ ,  $N_e$  is the electron density of the medium and  $G_F$  is Fermi's  
270 constant. It is clear to see that there exists a value of  $A = \Delta m^2 \cos(2\theta)$  for  $\Delta m^2 > 0$   
271 which results in a divergent mixing parameter. This resonance is termed the Mikheyev-  
272 Smirnov-Wolfenstein (MSW) effect (or more colloquially, the matter resonance) which  
273 regenerates the electron neutrino component of the neutrino flux [18–20]. The density  
274 at which the resonance occurs is given by

$$N_e = \frac{\Delta m^2 \cos(2\theta)}{2\sqrt{2}G_F E}. \quad (2.12)$$

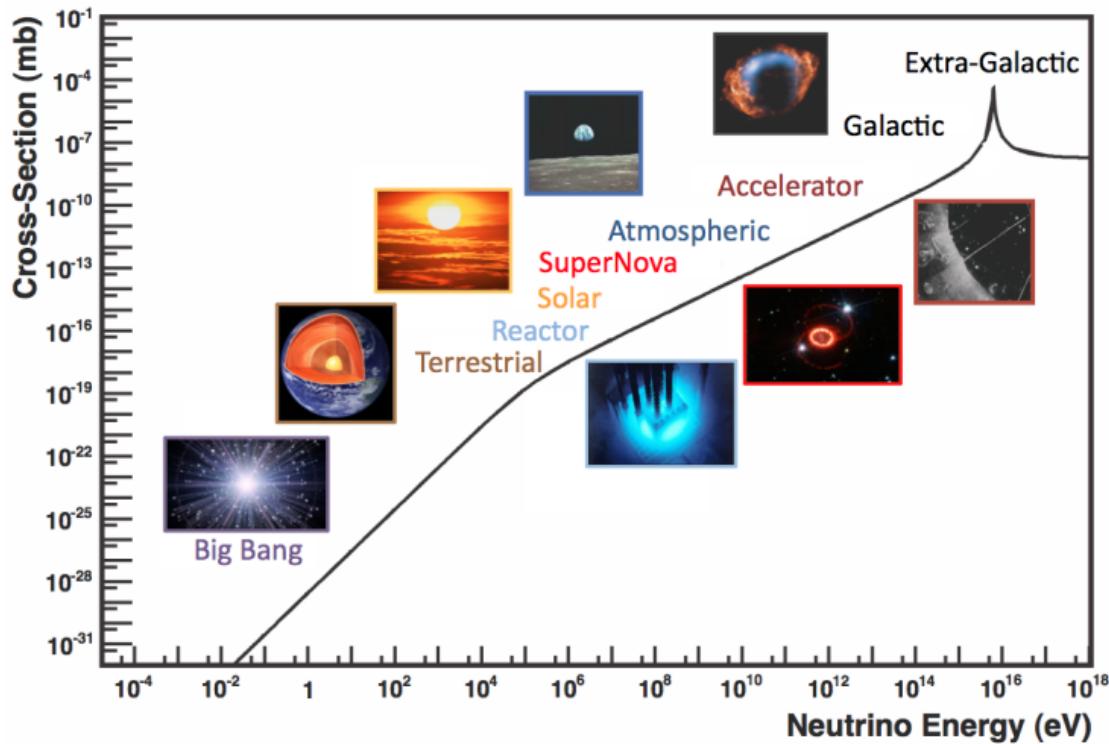
275 At densities lower than this critical value, the oscillation probability will be much  
276 closer to that of vacuum oscillation. For antineutrinos,  $N_e \rightarrow -N_e$  [21]. The resonance  
277 occurring from the MSW effect depends on the sign of  $\Delta m^2$ . Therefore, any neutrino  
278 oscillation experiment which observes neutrinos and antineutrinos which have propa-  
279 gated through matter can have some sensitivity to the ordering of the neutrino mass  
280 eigenstates.

## 281 2.3. Neutrino Oscillation Measurements

282 As evidence of beyond standard model physics, the 2015 Nobel Prize in Physics was  
283 awarded to the Super-Kamiokande (SK) [22] and Sudbury Neutrino Observatory  
284 (SNO) [23] collaborations for the first definitive observation of solar and atmospheric

<sup>285</sup> neutrino oscillation [24]. Since then, the field has seen a wide array of oscillation  
<sup>286</sup> measurements from a variety of neutrino sources. As seen in subsection 2.2.1, the  
<sup>287</sup> neutrino oscillation probability is dependent on the ratio of the propagation baseline,  $L$ ,  
<sup>288</sup> to the neutrino energy,  $E$ . It is this ratio that determines the type of neutrino oscillation  
<sup>289</sup> a particular experiment is sensitive to.

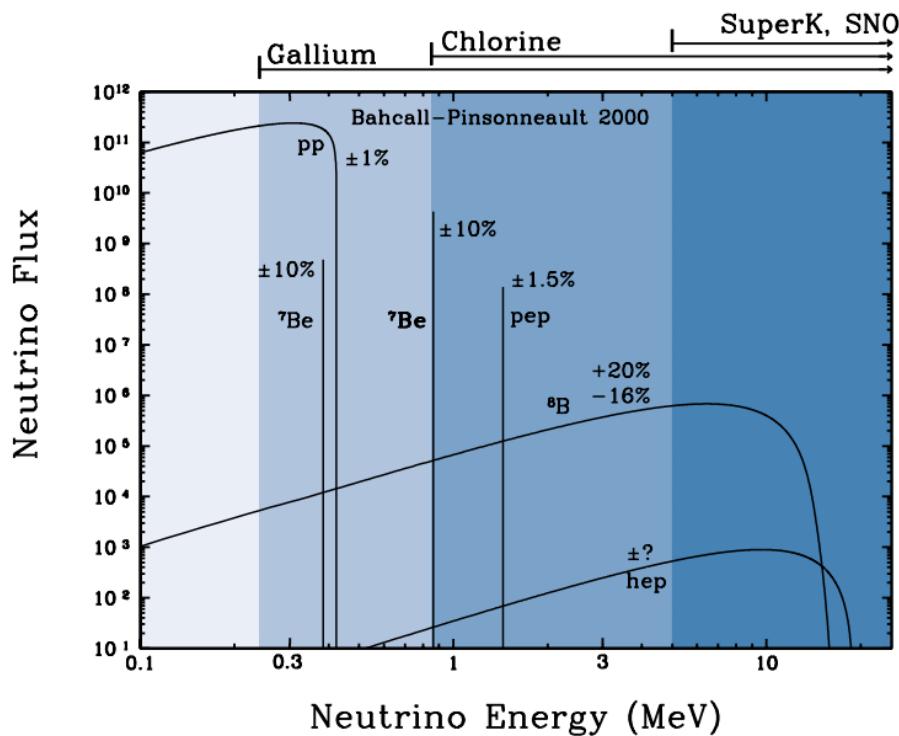
<sup>290</sup> As illustrated in Figure 2.1, there are many neutrino sources that span a wide  
<sup>291</sup> range of energies. The least energetic neutrinos are from diffuse supernovae and  
<sup>292</sup> terrestrial neutrinos at  $O(1)$  MeV whereas the most energetic neutrinos originate from  
<sup>293</sup> atmospheric and galactic neutrinos of  $> O(1)$  TeV.



**Figure 2.1.:** The cross-section of neutrinos from various natural and man-made sources as a function of neutrino energy. Taken from [25]

### <sup>294</sup> 2.3.1. Solar Neutrinos

<sup>295</sup> Solar neutrinos are emitted from fusion reaction chains at the center of the Sun. The  
<sup>296</sup> solar neutrino flux, given as a function of neutrino energy for different fusion and  
<sup>297</sup> decay chains, is illustrated in Figure 2.2. Whilst proton-proton fusion generates the  
<sup>298</sup> largest flux of neutrinos, the neutrinos are of low energy and are difficult to reconstruct  
<sup>299</sup> due to the IBD interaction threshold of 1.8MeV. Consequently, most experiments focus  
<sup>300</sup> on the neutrinos from the decay of  $^8B$  (via  $^8B \rightarrow ^8Be^* + e^+ + \nu_e$ ), which are higher  
<sup>301</sup> energy.

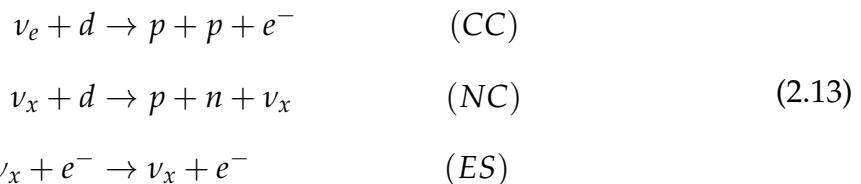


**Figure 2.2.:** The solar neutrino flux as a function of neutrino energy for various fusion reactions and decay chains as predicted by the Standard Solar Model. Taken from [26].

<sup>302</sup> The first measurements of solar neutrinos observed a significant reduction in the  
<sup>303</sup> event rate compared to predictions from the Standard Solar Model [27, 28]. The  
<sup>304</sup> proposed solution to this “solar neutrino problem” was  $\nu_e \leftrightarrow \nu_\mu$  oscillations in a

<sup>305</sup> precursory version of the PMNS model [29]. The Kamiokande [30], Gallex [31] and  
<sup>306</sup> Sage [32] experiments confirmed the  $\sim 0.5$  factor deficit of solar neutrinos.

<sup>307</sup> The conclusive solution to this problem was determined by the SNO collaboration  
<sup>308</sup> [33]. Using a deuterium water target to observe  ${}^8B$  neutrinos, the event rate of charged  
<sup>309</sup> current (CC), neutral current (NC), and elastic scattering (ES) interactions (Given in  
<sup>310</sup> Equation 2.13) was simultaneously measured. CC events can only occur for electron  
<sup>311</sup> neutrinos, whereas the NC channel is agnostic to neutrino flavour, and the ES reaction  
<sup>312</sup> has a slight excess sensitivity to electron neutrino interactions. This meant that there  
<sup>313</sup> were direct measurements of the  $\nu_e$  and  $\nu_x$  neutrino flux. It was concluded that the  
<sup>314</sup> CC and ES interaction rates were consistent with the deficit previously observed.  
<sup>315</sup> Most importantly, the NC reaction rate was only consistent with the others under the  
<sup>316</sup> hypothesis of flavour transformation.



<sup>317</sup> Many experiments have since measured the neutrino flux of different interaction  
<sup>318</sup> chains within the sun [34–36]. The most recent measurement was that of CNO neutrinos  
<sup>319</sup> which were recently observed with  $5\sigma$  significance by the Borexino collaboration.  
<sup>320</sup> Future neutrino experiments aim to further these spectroscopic measurements of  
<sup>321</sup> different fusion chains within the Sun [37–39]. Solar neutrinos act as an irreducible  
<sup>322</sup> background for dark matter experiments like DARWIN but oscillation parameter  
<sup>323</sup> measurements can be made [40].

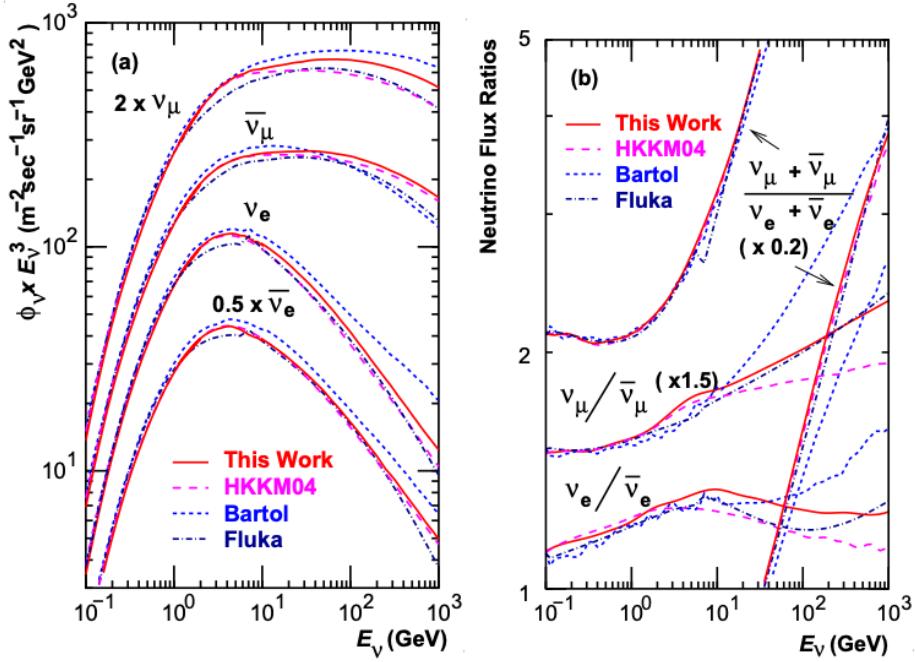
### <sup>324</sup> 2.3.2. Atmospheric Neutrinos

- <sup>325</sup> The interactions of primary cosmic ray protons in Earth's upper atmosphere generate  
<sup>326</sup> showers of energetic hadrons. These are mostly pions and kaons which when they  
<sup>327</sup> decay produce a natural source of neutrinos spanning energies of MeV to TeV [41].  
<sup>328</sup> The main decay is via

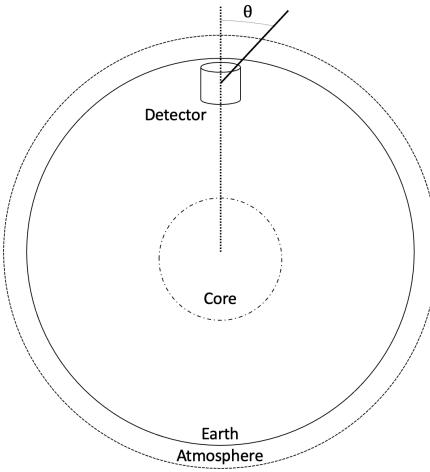
$$\begin{aligned} \pi^\pm &\rightarrow \mu^\pm + (\nu_\mu, \bar{\nu}_\mu) \\ \mu^\pm &\rightarrow e^\pm + (\nu_e, \bar{\nu}_e) + (\nu_\mu, \bar{\nu}_\mu) \end{aligned} \tag{2.14}$$

<sup>329</sup> such that for a single pion decay, three neutrinos are typically produced. The  
<sup>330</sup> atmospheric neutrino flux energy spectra as predicted by the Bartol [42], Honda  
<sup>331</sup> [43–45], and FLUKA [46] models are illustrated in Figure 2.3. The flux distribution  
<sup>332</sup> peaks at an energy of  $O(10)\text{GeV}$ . The uncertainties associated with these models  
<sup>333</sup> are dominated by the hadronic production of kaon and pions as well as the primary  
<sup>334</sup> cosmic flux.

<sup>335</sup> Unlike long-baseline experiments which have a fixed baseline, the distance at-  
<sup>336</sup> mospheric neutrinos propagate is dependent upon the zenith angle at which they  
<sup>337</sup> interact. This is illustrated in Figure 2.4. Neutrinos that are generated directly above  
<sup>338</sup> the detector ( $\cos(\theta) = 1.0$ ) have a baseline equivalent to the height of the atmosphere  
<sup>339</sup> whereas neutrinos that interact directly below the detector ( $\cos(\theta) = -1.0$ ) have to  
<sup>340</sup> travel a length equal to the diameter of the Earth. This means atmospheric neutrinos  
<sup>341</sup> have a baseline that varies from  $O(20)\text{km}$  to  $O(6 \times 10^3)\text{km}$ . Any neutrino generated  
<sup>342</sup> at or below the horizon will be subject to matter effects as they propagate through the  
<sup>343</sup> Earth.



**Figure 2.3.:** Left panel: The atmospheric neutrino flux for different neutrino flavours as a function of neutrino energy as predicted by the 2007 Honda model (“This work”) [43], the 2004 Honda model (“HKKM04”) [44], the Bartol model [42] and the FLUKA model [46]. Right panel: The ratio of the muon to electron neutrino flux as predicted by all the quoted models. Both figures taken from [43].



**Figure 2.4.:** A diagram illustrating the definition of zenith angle as used in the Super Kamiokande experiment [47].

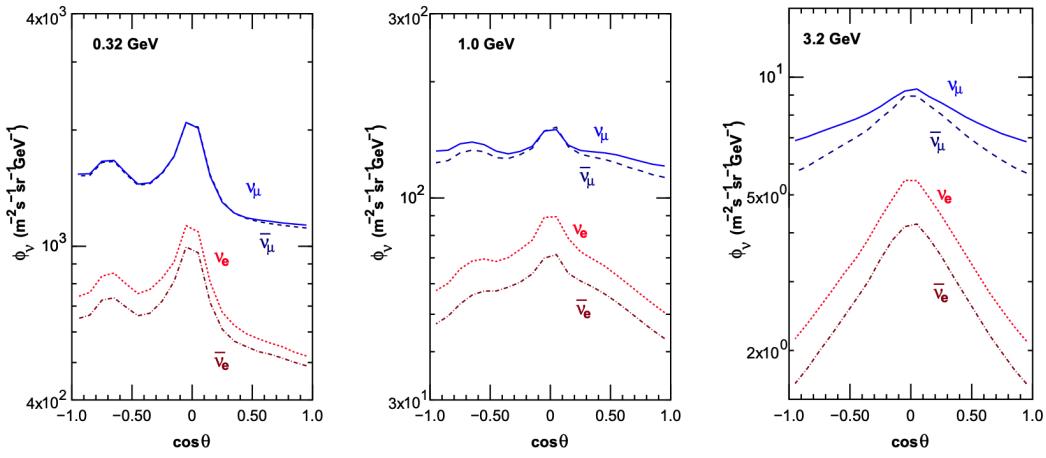
Figure 2.5 highlights the neutrino flux as a function of the zenith angle for different

slices of neutrino energy. For medium to high-energy neutrinos (and to a lesser degree

for low-energy neutrinos), the flux is approximately symmetric around  $\cos(\theta) = 0$ .

<sup>344</sup> Figure 2.5 highlights the neutrino flux as a function of the zenith angle for different  
<sup>345</sup> slices of neutrino energy. For medium to high-energy neutrinos (and to a lesser degree  
<sup>346</sup> for low-energy neutrinos), the flux is approximately symmetric around  $\cos(\theta) = 0$ .

<sup>347</sup> To the accuracy of this approximation, the systematic uncertainties associated with  
<sup>348</sup> atmospheric flux for comparing upward-going and down-going neutrino cancels. This  
<sup>349</sup> allows the down-going events, which are mostly insensitive to oscillation probabilities,  
<sup>350</sup> to act as an unoscillated prediction (similar to a near detector in an accelerator neutrino  
<sup>351</sup> experiment).

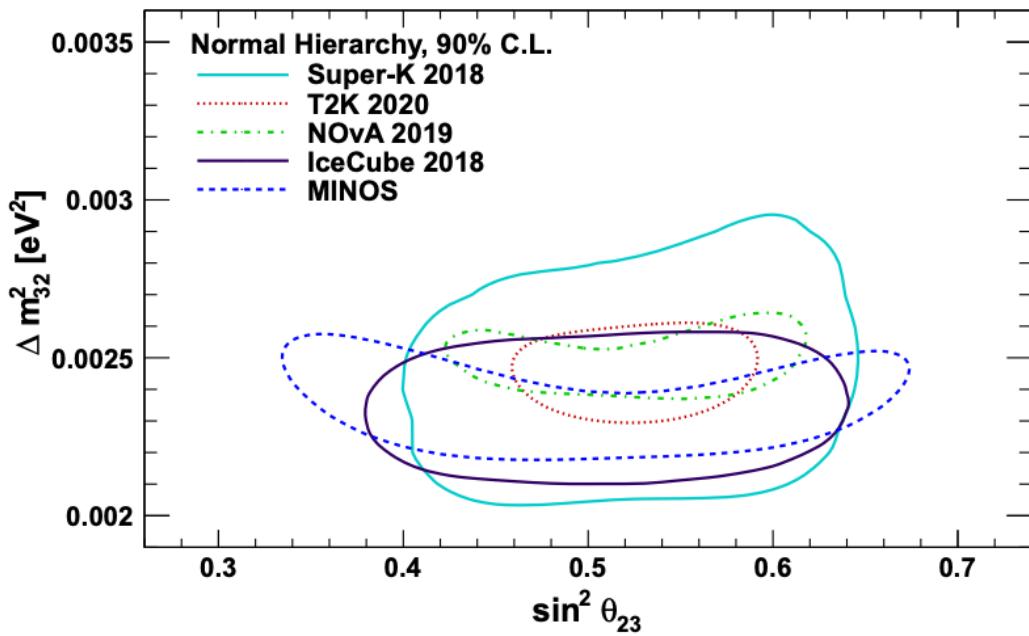


**Figure 2.5.:** Prediction of  $\nu_e$ ,  $\bar{\nu}_e$ ,  $\nu_\mu$ ,  $\bar{\nu}_\mu$  fluxes as a function of zenith angle as calculated by the HKKM model [45]. The left, middle and right panels represent three values of neutrino energy, 0.32GeV, 1.0GeV and 3.2GeV respectively. Predictions for other models including Bartol [42], Honda [43] and FLUKA [46] are given in [47].

<sup>352</sup> Precursory hints of atmospheric neutrinos were observed in the mid-1960s search-  
<sup>353</sup> ing for  $\nu_\mu + X \xrightarrow{(-)} X^* + \mu^\pm$  [48], although it was called an anomaly at the time of  
<sup>354</sup> measurement. This was succeeded with the IMB-3 [49] and Kamiokande [50] experi-  
<sup>355</sup> ments which measured the ratio of muon neutrinos compared to electron neutrinos  
<sup>356</sup>  $R(\nu_\mu/\nu_e)$ . Both experiments were found to have a consistent deficit of muon neutrinos,  
<sup>357</sup> with  $R(\nu_\mu/\nu_e) = 0.67 \pm 0.17$  and  $R(\nu_\mu/\nu_e) = 0.60^{+0.07}_{-0.06} \pm 0.05$ . Super-Kamiokande  
<sup>358</sup> (SK) [47] extended this analysis by fitting oscillation parameters in  $P(\nu_\mu \rightarrow \nu_\tau)$  which  
<sup>359</sup> found best fit parameters  $\sin^2(2\theta) > 0.92$  and  $1.5 \times 10^{-3} < \Delta m^2 < 3.4 \times 10^{-3} \text{ eV}^2$ .

<sup>360</sup> Since then, atmospheric neutrino experiments have been making precision mea-  
<sup>361</sup> surements of the  $\sin^2(\theta_{23})$  and  $\Delta m^2_{32}$  oscillation parameters. Atmospheric neutrino  
<sup>362</sup> oscillation is dominated by  $P(\nu_\mu \rightarrow \nu_\tau)$ , where SK observed a  $4.6\sigma$  discovery of  $\nu_\tau$

<sup>363</sup> appearance [51]. Figure 2.6 illustrates the current estimates on the atmospheric mixing  
<sup>364</sup> parameters from a wide range of atmospheric and accelerator neutrino observatories.



**Figure 2.6.:** Constraints on the atmospheric oscillation parameters,  $\sin^2(\theta_{23})$  and  $\Delta m_{23}^2$ , from atmospheric and long baseline experiments: SK [52], T2K [53], NOvA [54], IceCube [55] and MINOS [56]. Figure taken from [57].

### <sup>365</sup> 2.3.3. Accelerator Neutrinos

<sup>366</sup> The concept of using a man-made “neutrino beam” was first realised in 1962 [58].  
<sup>367</sup> Since then, many experiments have followed which all use the same fundamental  
<sup>368</sup> concepts. Typically, a proton beam is aimed at a target producing charged mesons that  
<sup>369</sup> decay to neutrinos. The mesons can be sign-selected by the use of magnetic focusing  
<sup>370</sup> horns to generate a neutrino or antineutrino beam. Pions are the primary meson that  
<sup>371</sup> decay and depending on the orientation of the magnetic field, a muon (anti-)neutrino  
<sup>372</sup> beam is generated via  $\pi^+ \rightarrow \mu^+ + \nu_\mu$  or  $\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$ . The decay of muons and  
<sup>373</sup> kaons does result in an irreducible intrinsic electron neutrino background. In T2K,  
<sup>374</sup> this background contamination is  $O(< 1\%)$  [59]. There is also an approximately  $\sim 5\%$

<sup>375</sup> “wrong-sign” neutrino background of  $\bar{\nu}_\mu$  generated via the same decays. As the beam is  
<sup>376</sup> generated by proton interactions (rather than anti-proton interactions), the wrong-sign  
<sup>377</sup> component in the antineutrino beam is larger when operating in neutrino mode.

<sup>378</sup> Tuning the proton energy in the beam and using beam focusing techniques allows  
<sup>379</sup> the neutrino energy to be set to a value that maximises the disappearance oscillation  
<sup>380</sup> probability in the  $L/E$  term in Equation 2.10. This means that accelerator experiments  
<sup>381</sup> are typically more sensitive to the mixing parameters as compared to a natural neutrino  
<sup>382</sup> source. However, the disadvantage compared to atmospheric neutrino experiments is  
<sup>383</sup> that the baseline has to be shorter due to the lower flux. Consequently, there is typically  
<sup>384</sup> less sensitivity to matter effects and the ordering of the neutrino mass eigenstates.

<sup>385</sup> A neutrino experiment measures

$$R(\vec{x}) = \Phi(E_\nu) \times \sigma(E_\nu) \times \epsilon(\vec{x}) \times P(\nu_\alpha \rightarrow \nu_\beta), \quad (2.15)$$

<sup>386</sup> where  $R(\vec{x})$  is the event rate of neutrinos at position  $\vec{x}$ ,  $\Phi(E_\nu)$  is the flux of neutrinos  
<sup>387</sup> with energy  $E_\nu$ ,  $\sigma(E_\nu)$  is the cross-section of the neutrino interaction and  $\epsilon(\vec{x})$  is the  
<sup>388</sup> efficiency and resolution of the detector. In order to leverage the most out of an  
<sup>389</sup> accelerator neutrino experiment, the flux and cross-section systematics need to be  
<sup>390</sup> constrained. This is typically done via the use of a “near detector”, situated at a baseline  
<sup>391</sup> of  $O(1)$ km. This detector observes the unoscillated neutrino flux and constrains the  
<sup>392</sup> parameters used within the flux and cross-section model.

<sup>393</sup> The first accelerator experiments to precisely measure oscillation parameters were  
<sup>394</sup> MINOS [60] and K2K [61]. These experiments confirmed the  $\nu_\mu$  disappearance seen in  
<sup>395</sup> atmospheric neutrino experiments by finding consistent parameter values for  $\sin^2(\theta_{23})$   
<sup>396</sup> and  $\Delta m_{23}^2$ . The current generation of accelerator neutrino experiments, T2K and NO $\nu$ A

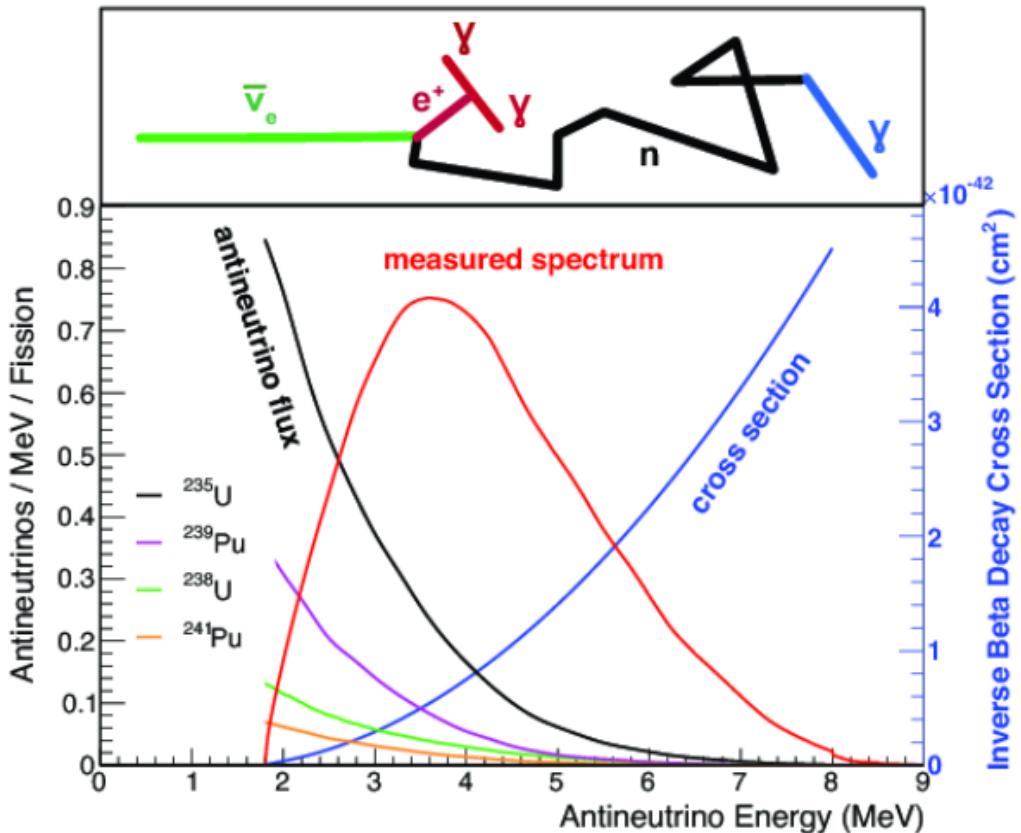
<sup>397</sup> extended this field by observing  $\bar{\nu}_\mu \rightarrow \bar{\nu}_e$  and lead the sensitivity to atmospheric mix-  
<sup>398</sup> ing parameters as seen in Figure 2.6 [62]. The two experiments differ in their peak  
<sup>399</sup> neutrino energy, baseline, and detection technique. The NO $\nu$ A experiment is situated  
<sup>400</sup> at a baseline of 810km from the NuMI beamline which delivers 2GeV neutrinos. The  
<sup>401</sup> T2K neutrino beam is peaked around 0.6GeV and propagates 295km. The NO $\nu$ A  
<sup>402</sup> experiment also uses functionally identical detectors (near and far) which allow the  
<sup>403</sup> approximate cancellation of detector systematics whereas T2K uses a plastic scintil-  
<sup>404</sup> lator technique at the near detector and a water Cherenkov far detector. The future  
<sup>405</sup> generation experiments DUNE [63] and Hyper-Kamiokande [64] will succeed these  
<sup>406</sup> experiments as the high-precision era of neutrino oscillation parameter measurements  
<sup>407</sup> develops.

<sup>408</sup> Several anomalous results have been observed in the LSND [9] and MiniBooNE [10]  
<sup>409</sup> detectors which were designed with purposefully short baselines. Parts of the neu-  
<sup>410</sup> trino community attributed these results to oscillations induced by a fourth “sterile”  
<sup>411</sup> neutrino [65] but several searches in other experiments, MicroBooNE [66] and KAR-  
<sup>412</sup> MEN [67], found no hints of additional neutrino species. The solution to the anomalous  
<sup>413</sup> results is still being determined.

#### <sup>414</sup> 2.3.4. Reactor Neutrinos

<sup>415</sup> As illustrated in the first discovery of neutrinos (section 2.1), nuclear reactors are a very  
<sup>416</sup> useful man-made source of electron antineutrinos. For reactors that use low-enriched  
<sup>417</sup> uranium  $^{235}\text{U}$  as fuel, the antineutrino flux is dominated by the  $\beta$ -decay fission of  $^{235}\text{U}$ ,  
<sup>418</sup>  $^{238}\text{U}$ ,  $^{239}\text{Pu}$  and  $^{241}\text{Pu}$  [68] as illustrated in Figure 2.7.

<sup>419</sup> Due to their low energy, reactor electron antineutrinos predominantly interact  
<sup>420</sup> via the inverse  $\beta$ -decay (IBD) interaction. The typical signature contains two signals



**Figure 2.7.:** Reactor electron antineutrino fluxes for  $^{235}\text{U}$  (Black),  $^{238}\text{U}$  (Green),  $^{239}\text{Pu}$  (Purple), and  $^{241}\text{Pu}$  (Orange) isotopes. The inverse  $\beta$ -decay cross-section (Blue) and corresponding measurable neutrino spectrum (Red) are also given. Top panel: Schematic of Inverse  $\beta$ -decay interaction including the eventual capture of the emitted neutron. This capture emits a  $\gamma$ -ray which provides a second signal of the event. Taken from [69].

delayed by  $O(200)\mu\text{s}$ ; firstly the prompt photons from positron annihilation, and secondly the photons emitted ( $E_{tot}^\gamma = 2.2\text{MeV}$ ) from de-excitation after neutron capture on hydrogen. Searching for both signals improves the detector's ability to distinguish between background and signal events [70]. Recently, SK included gadolinium dopants into the ultra-pure water to increase the energy released from the photon cascade to  $\sim 8\text{MeV}$  and reduce the time of the delayed signal to  $\sim 28\mu\text{s}$ .

There are many short baseline experiments ( $L \sim O(1)\text{km}$ ) that have measured the  $\sin^2(\theta_{13})$  and  $\Delta m_{23}^2$  oscillation parameters. Daya Bay [71], RENO [72] and Double Chooz [73] have all provided precise measurements, with the first discovery of a

430 non-zero  $\theta_{13}$  made by Daya Bay and RENO (and complemented by T2K [73]). The  
431 constraints on  $\sin^2(\theta_{13})$  by the reactor experiments lead the field and are often used as  
432 external inputs to accelerator neutrino experiments to improve their sensitivity to  $\delta_{CP}$   
433 and mass hierarchy determination. JUNO-TAO [74], a small collaboration within the  
434 larger JUNO experiment, is a next-generation reactor experiment that aims to precisely  
435 measure the isotopic antineutrino yields from the different fission chains. Alongside  
436 this, it aims to explain the ‘5MeV excess’ [75–77] by conducting a search for sterile  
437 neutrinos with a mass scale of around 1eV.

438 Kamland [78] is the only experiment to have observed reactor neutrinos using a  
439 long baseline (flux weighted averaged baseline of  $L \sim 180\text{km}$ ) which allows it to have  
440 sensitivity to  $\Delta m_{12}^2$ . Combined with the SK solar neutrino experiment, the combined  
441 analysis puts the most stringent constraint on  $\Delta m_{12}^2$  [79].

## 442 2.4. Summary

443 Since observing the first evidence of neutrino oscillations in the late 1990’s, numerous  
444 measurements of the mixing parameters have been made. Many experiments use  
445 neutrinos as a tool for discovery of new physics (diffuse supernova background,  
446 neutrinoless double beta decay and others) so the PMNS parameters are summarised  
447 in the Particle Data Group (PDG) review tables. The analysis presented in this thesis  
448 focuses on the 2020 T2K oscillation analysis presented in [80] where the 2018 PDG  
449 constraints [81] were used. These constraints are outlined in Table 2.1.

450 The  $\sin^2(\theta_{13})$  measurement stems from the electron antineutrino disappearance,  
451  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$ , and is take as the average best-fit from the combination of Daya Bay,  
452 Reno and Double Chooz. It is often used as a prior uncertainty within other neu-  
453 trino oscillation experiments, typically termed the reactor constraint. The  $\sin^2(\theta_{12})$

Parameter	2018 Constraint
$\sin^2(\theta_{12})$	$0.307 \pm 0.013$
$\Delta m_{21}^2$	$(7.53 \pm 0.18) \times 10^{-5} \text{ eV}^2$
$\sin^2(\theta_{13})$	$(2.12 \pm 0.08) \times 10^{-2}$
$\sin^2(\theta_{23})$ (I.H., Q1)	$0.421^{+0.033}_{-0.025}$
$\sin^2(\theta_{23})$ (I.H., Q2)	$0.592^{+0.023}_{-0.030}$
$\sin^2(\theta_{23})$ (N.H., Q1)	$0.417^{+0.025}_{-0.028}$
$\sin^2(\theta_{23})$ (N.H., Q2)	$0.597^{+0.024}_{-0.030}$
$\Delta m_{32}^2$ (I.H.)	$(-2.56 \pm 0.04) \times 10^{-3} \text{ eV}^2$
$\Delta m_{32}^2$ (N.H.)	$(2.51 \pm 0.05) \times 10^{-3} \text{ eV}^2$

**Table 2.1.:** The 2018 Particle Data Group constraints of the oscillation parameters taken from [81]. The value of  $\Delta m_{23}^2$  is given for both normal hierarchy (N.H.) and inverted hierarchy (I.H.) and  $\sin^2(\theta_{23})$  is broken down by whether its value is below (Q1) or above (Q2) 0.5.

parameter is predominantly measured through electron neutrino disappearance,  $P(\nu_e \rightarrow \nu_{\mu,\tau})$ , in solar neutrino experiments. The long-baseline reactor neutrino experiment Kamland also has sensitivity to this parameter and is used in a joint fit to solar data from SNO and SK, using the reactor constraint. Measurements of  $\sin^2(\theta_{23})$  are made by long-baseline and atmospheric neutrino experiments. The PDG value is a joint fit of T2K, NOvA, MINOS and IceCube DeepCore experiments. The latest T2K-only measurement, provided at Neutrino2020 and is the basis of this thesis, is given as  $\sin^2(\theta_{23}) = 0.546^{+0.024}_{-0.046}$  [80]. The PDG constraint on  $\Delta m_{12}^2$  is provided by the KamLAND experiment using solar and geoneutrino data. This measurement utilised a  $\sin^2(\theta_{13})$  constraint from accelerator (T2K, MINOS) and reactor neutrino (Daya Bay, RENO, Double Chooz) experiments. Accelerator measurements make some of the most stringent constraints on  $\Delta m_{23}^2$  although atmospheric experiments have more sensitivity to the mass hierarchy determination. The PDG performs a joint fit of accelerator and atmospheric data, in both normal and inverted hierarchy separately. The latest T2K-only result is  $\Delta m_{32}^2 = 2.49^{+0.058}_{-0.082} \times 10^{-3} \text{ eV}^2$  favouring normal hierarchy [80]. The value of  $\delta_{CP}$  is largely undetermined. CP-conserving values of 0 and  $\pi$  were

<sup>470</sup> rejected with  $\sim 2\sigma$  intervals, as published in Nature, although more recent analysis  
<sup>471</sup> have reduced the rejection intervals to 90%. Since the 2018 PDG publication, there has  
<sup>472</sup> been a new measurement of  $\sin^2(\theta_{13}) = (2.20 \pm 0.07) \times 10^{-2}$  [82], alongside updated  
<sup>473</sup>  $\Delta m_{23}^2$  and  $\sin^2(\theta_{23})$  measurements.

<sup>474</sup> Throughout this thesis, several sample spectra predictions and contours are pre-  
<sup>475</sup> sented which require oscillation parameters to be assumed. Table 2.2 defines two sets  
<sup>476</sup> of oscillation parameters, with “Asimov A” set being close to the preferred values  
<sup>477</sup> from a previous T2K-only fit [83] and “Asimov B” being CP-conserving and further  
<sup>478</sup> from maximal  $\theta_{23}$  mixing.

Parameter	Asimov A	Asimov B
$\Delta m_{12}^2$	$7.53 \times 10^{-5} \text{ eV}^2$	
$\Delta m_{32}^2$	$2.509 \times 10^{-3} \text{ eV}^2$	
$\sin^2(\theta_{12})$	0.304	
$\sin^2(\theta_{13})$	0.0219	
$\sin^2(\theta_{23})$	0.528	0.45
$\delta_{CP}$	-1.601	0.0

**Table 2.2.:** Reference values of the neutrino oscillation parameters for two different oscillation parameter sets.

<sup>479</sup> **Chapter 3**

<sup>480</sup> **T2K and SK Experiment Overview**

<sup>481</sup> As the successor of the Kamiokande experiment, the Super-Kamiokande (SK) collabora-  
<sup>482</sup> ration has been leading atmospheric neutrino oscillation analyses for over two decades.  
<sup>483</sup> The detector has provided some of the strongest constraints on proton decay and the  
<sup>484</sup> first precise measurements of the  $\Delta m_{23}^2$  and  $\sin^2(\theta_{23})$  neutrino oscillation parameters.  
<sup>485</sup> The ability of the detector to low-energy neutrino events has been significantly im-  
<sup>486</sup> proved with the recent gadolinium doping of the ultra-pure water target. The history,  
<sup>487</sup> detection technique, and operation of the SK detector is described in section 3.1.

<sup>488</sup> The Tokai-to-Kamioka (T2K) experiment was one of the first long-baseline ex-  
<sup>489</sup> periments to use both neutrino and antineutrino beams to precisely measure the  
<sup>490</sup> charge parity violation within the neutrino sector. With the SK detector observing  
<sup>491</sup> the oscillated neutrino flux, the T2K experiment observed the first hints of a non-zero  
<sup>492</sup>  $\sin^2(\theta_{13})$  measurement and continues to lead the field with the constraints it provides  
<sup>493</sup> on  $\sin^2(\theta_{13})$ ,  $\sin^2(\theta_{23})$ ,  $\Delta m_{23}^2$  and  $\delta_{CP}$ . The techniques which T2K uses in gener-  
<sup>494</sup> ating its neutrino beam as well as the near-detector used to constrain the flux and  
<sup>495</sup> cross-section parameters used in this analysis are documented in section 3.2.

<sup>496</sup> **3.1. The Super-Kamiokande Experiment**

<sup>497</sup> The SK experiment began taking data in 1996 [84] and has had many modifications  
<sup>498</sup> throughout its lifespan. There have been seven defined periods of data taking as  
<sup>499</sup> noted in Table 3.1. Data taking began in SK-I which ran for five years. Between the

500 SK-I and SK-II periods, a significant proportion of the PMTs were damaged during  
 501 maintenance. Those that survived were equally distributed throughout the detector  
 502 in the SK-II era, which resulted in a reduced photo-coverage. From SK-III onwards,  
 503 repairs to the detector meant the full suite of PMTs was operational. Before the  
 504 start of SK-IV, the data acquisition and electronic systems were upgraded. Between  
 505 SK-IV and SK-V, a significant effort was placed into tank open maintenance and  
 506 repair/replacement of defective PMTs, a task for which the author of this thesis was  
 507 required. Consequently, the detector conditions were significantly different between  
 508 the two operational periods. SK-VI saw the start of the 0.01% gadolinium doped water.  
 509 SK-VII, which started during the writing of this thesis, has increased the gadolinium  
 510 concentration to 0.03% for continued operation [85].

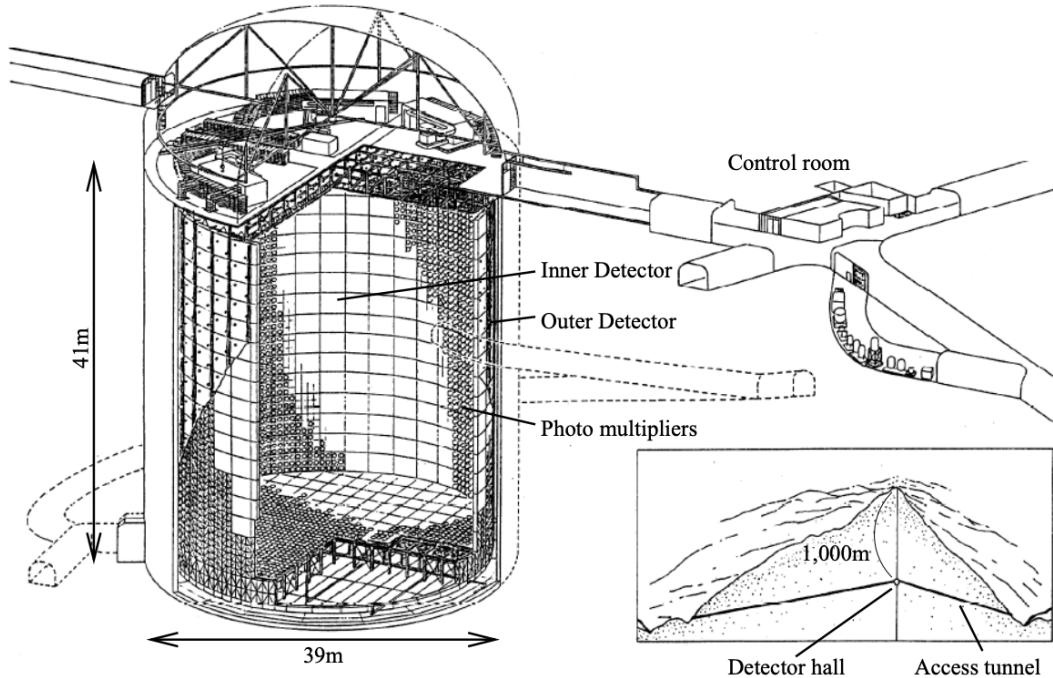
Period	Start Date	End Date	Live-time (days)
I	April 1996	July 2001	1489.19
II	October 2002	October 2005	798.59
III	July 2006	September 2008	518.08
IV	September 2008	May 2018	3244.4
V	January 2019	July 2020	461.02
VI	July 2020	May 2022	583.3
VII	May 2022	Ongoing	N/A

**Table 3.1.:** The various SK periods and respective live-time. The SK-VI live-time is calculated until 1<sup>st</sup> April 2022. SK-VII started during the writing of this thesis.

### 511 3.1.1. The SK Detector

512 The basic structure of the Super-Kamiokande (SK) detector is a cylindrical tank with a  
 513 diameter 39.3m and height 41.1m filled with ultrapure water [86]. A diagram of the  
 514 significant components of the SK detector is given in Figure 3.1. The SK detector is  
 515 situated in the Kamioka mine in Gifu, Japan. The mine is underground with roughly  
 516 1km rock overburden (2.7km water equivalent overburden) [87]. At this depth, the

<sup>517</sup> rate of cosmic ray muons is significantly decreased to a value of  $\sim 2\text{Hz}$ . The top of  
<sup>518</sup> the tank is covered with stainless steel which is designed as a working platform for  
<sup>519</sup> maintenance, calibration, and location for high voltage and data acquisition electronics.



**Figure 3.1.:** A schematic diagram of the Super-Kamiokande Detector. Taken from [88].

<sup>520</sup> A smaller cylindrical structure (36.2m diameter, 33.8m height) is situated inside the  
<sup>521</sup> tank, with an approximate 2m gap between this structure and the outer tank wall. The  
<sup>522</sup> purpose of this structure is to support the photomultiplier tubes (PMTs). The volume  
<sup>523</sup> inside and outside the support structure is referred to as the inner detector (ID) and  
<sup>524</sup> outer detector (OD), respectively. In the SK-IV era, the ID and OD are instrumented  
<sup>525</sup> by 11,129 50cm and 1,885 20cm PMTs respectively [86]. The ID contains a 32kton  
<sup>526</sup> mass of water. Many analyses performed at SK use a “fiducial volume” defined by the  
<sup>527</sup> volume of water inside the ID excluding some distance to the ID wall. This reduces the  
<sup>528</sup> volume of the detector which is sensitive to neutrino events but reduces radioactive  
<sup>529</sup> backgrounds and allows for better reconstruction performance. The nominal fiducial

volume is defined as the area contained inside 2m from the ID wall for a total of 22.5kton water [89].

The two regions of the detector (ID and OD) are optically separated with opaque black plastic. The purpose of this is to determine whether a track entered or exited the ID. This allows cosmic ray muons and partially contained events to be tagged and separated from neutrino events entirely contained within the ID. This black plastic is also used to cover the area between the ID PMTs to reduce photon reflection from the ID walls. Opposite to this, the OD is lined with a reflective material to allow photons to reflect around inside the OD until collected by one of the PMTs. Furthermore, each OD PMT is backed with  $50 \times 50\text{cm}$  plates of wavelength shifting acrylic which increases the efficiency of light collection [87].

In the SK-IV data-taking period, the photocathode coverage of the detector, or the fraction of the ID wall instrumented with PMTs, is  $\sim 40\%$  [87]. The PMTs have a quantum efficiency (the ratio of detected electrons to incident photons) of  $\sim 21\%$  for photons with wavelengths of  $360\text{nm} < \lambda < 390\text{nm}$ . The proportion of photoelectrons that produce a signal in the dynode of a PMT, termed the collection efficiency, is  $> 70\%$  [87]. The PMTs used within SK are most sensitive to photons with wavelength  $300\text{nm} \leq \lambda \leq 600\text{nm}$  [87]. One disadvantage of using PMTs as the detection media is that the Earth's geomagnetic field can modify its response. Therefore, a set of compensation coils is built around the inner surface of the detector to mitigate this effect [90].

As mentioned, the SK detector is filled with ultrapure water, which in a perfect world would contain no impurities. However, bacteria and organic compounds can significantly degrade the water quality. This decreases the attenuation length, which reduces the total number of photons that hit a PMT. To combat this, a sophisticated water treatment system has been developed [87, 91]. UV lights, mechanical filters,

and membrane degasifiers are used to reduce the bacteria, suspended particulates, and radioactive materials from the water. The flow of water within the tank is also critical as it can remove stagnant bacterial growth or build-up of dust on the surfaces within the tank. Gravity drifts impurities in the water towards the bottom of the tank which, if left uncontrolled, can create asymmetric water conditions between the top and bottom of the tank. Typically, the water entering the tank is cooled below the ambient temperature of the tank to control convection and inhibit bacteria growth. Furthermore, the rate of dark noise hits within PMTs is sensitive to the PMT temperature [92] so controlling the temperature gradients within the tank is beneficial for stable measurements.

SK-VI is the first phase of the SK experiment to use gadolinium dopants within the ultrapure water [85]. As such, the SK water system had to be replaced to avoid removing the gadolinium concentrate from the ultrapure water [93]. For an inverse  $\beta$ -decay (IBD) interaction in a water target, the emitted neutron is thermally captured on hydrogen. This process releases 2.2MeV  $\gamma$  rays which are difficult to detect as the resulting Compton scattered electrons are very close to the Cherenkov threshold, limiting the number of photons produced. Thermal capture of neutrons on gadolinium generates  $\gamma$  rays with higher energy (8MeV [70]) meaning they are more easily detected. SK-VI has 0.01% Gd loading (0.02% gadolinium sulphate by mass) which causes  $\approx$  50% of neutrons emitted by IBD to be captured on gadolinium [94, 95]. Whilst predominantly useful for low energy analyses, Gd loading allows better  $\nu/\bar{\nu}$  separation for atmospheric neutrino event selections [96]. Efforts are currently in place to increase the gadolinium concentrate to 0.03% for  $\approx$  75% neutron capture efficiency on gadolinium [97]. The final stage of loading targets 0.1% concentrate.

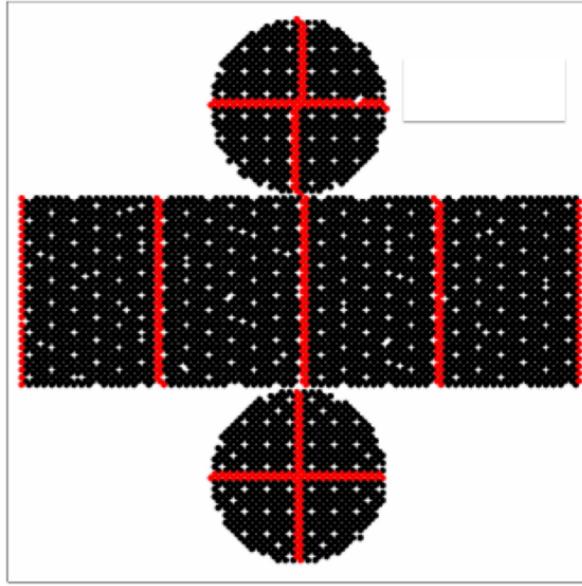
### 580 3.1.2. Calibration

581 The calibration of the SK detector is documented in [86] and summarised below. The  
582 analysis presented within this thesis is dependent upon ‘high energy events’ (Charged  
583 particles with  $O(> 100)\text{MeV}$  momenta). These are events that are expected to generate  
584 a larger number of photons such that each PMT will be hit with multiple photons.  
585 The reconstruction of these events depends upon the charge deposited within each  
586 PMT and the timing response of each individual PMT. Therefore, the most relevant  
587 calibration techniques to this thesis are outlined.

588 Before installation, 420 PMTs were calibrated to have identical charge responses  
589 and then distributed throughout the tank in a cross-shape pattern (As illustrated by  
590 Figure 3.2). These are used as a standardised measure for the rest of the PMTs installed  
591 at similar geometric positions within SK to be calibrated against. To perform this  
592 calibration, a xenon lamp is located at the center of the SK tank which flashes uniform  
593 light at 1Hz. This allows for geometrical effects, water quality variation, and timing  
594 effects to be measured in-situ throughout normal data-taking periods.

595 When specifically performing calibration of the detector (in out-of-data taking  
596 mode), the water in the tank was circulated to avoid top/bottom asymmetric water  
597 quality. Any non-uniformity within the tank significantly affects the PMT hit proba-  
598 bility through scattering or absorption. This becomes a dominant effect for the very  
599 low-intensity light sources discussed later which are designed such that only one  
600 photon is incident upon a given PMT.

601 The “gain” of a PMT is defined as the ratio of the total charge of the signal produced  
602 compared to the charge of photoelectrons emitted by the photocathodes within the  
603 PMT. To calibrate the signal of each PMT, the “relative” and “absolute” gain values are



**Figure 3.2.:** The location of “standard PMTs” (red) inside the SK detector. Taken from [86].

604 measured. The relative gain is the variation of gain among each of the PMTs whereas  
 605 the absolute gain is the average gain of all PMTs.

606 The relative gain is calibrated as follows. A laser is used to generate two measure-  
 607 ments: a high-intensity flash that illuminates every PMT with a sufficient number of  
 608 photons, and a low-intensity flash in which only a small number of PMTs collect light.  
 609 The first measurement creates an average charge,  $Q_{obs}(i)$  on PMT  $i$ , whereas the second  
 610 measurement ensures that each hit PMT only generates a single photoelectron. For the  
 611 low-intensity measurement, the number of times each PMT records a charge larger  
 612 than 1/4 photoelectrons,  $N_{obs}(i)$ , is counted. The values measured can be expressed as

$$\begin{aligned} Q_{obs}(i) &\propto I_H \times f(i) \times \epsilon(i) \times G(i), \\ N_{obs}(i) &\propto I_L \times f(i) \times \epsilon(i), \end{aligned} \tag{3.1}$$

613 Where  $I_H$  and  $I_L$  is the intensity of the high and low flashes,  $f(i)$  is the acceptance  
 614 efficiency of the  $i^{\text{th}}$  PMT,  $\epsilon(i)$  is the product of the quantum and collection efficiency

615 of the  $i^{\text{th}}$  PMT and  $G(i)$  is the gain of the  $i^{\text{th}}$  PMT. The relative gain for each PMT can  
616 determined by taking the ratio of these quantities.

617 The absolute gain calibration is performed by observing fixed energy  $\gamma$ -rays of  
618  $E_{\gamma} \sim 9\text{MeV}$  emitted isotropically from neutron capture on a NiCf source situated at  
619 the center of the detector. This generates a photon yield of about 0.004 photoelec-  
620 trons/PMT/event, meaning that  $> 99\%$  of PMT signals are generated from single  
621 photoelectrons. A charge distribution is generated by performing this calibration over  
622 all PMTs, and the average value of this distribution is taken to be the absolute gain  
623 value.

624 As mentioned in subsection 3.1.1, the average quantum and collection efficiency  
625 for the SK detector is  $\sim 21\%$  and  $> 70\%$  respectively. However, these values do differ  
626 between each PMT and need to be calibrated accordingly. Consequently, the NiCf  
627 source is also used to calibrate the “quantum  $\times$  collection” efficiency (denoted “QE”)  
628 value of each PMT. The NiCf low-intensity source is used as the PMT hit probability  
629 is proportional to the QE ( $N_{\text{obs}}(i) \propto \epsilon(i)$  in Equation 3.1). A Monte Carlo prediction  
630 which includes photon absorption, scattering, and reflection is made to estimate the  
631 number of photons incident on each PMT and the ratio of the number of predicted  
632 to observed hits is calculated. The difference is attributed to the QE efficiency of that  
633 PMT. This technique is extended to calculate the relative QE efficiency by normalizing  
634 the average of all PMTs which removes the dependence on the light intensity.

635 Due to differing cable lengths and readout electronics, the timing response between  
636 a photon hitting the PMT and the signal being captured by the data acquisition can be  
637 different between each PMT. Due to threshold triggers (Described in subsection 3.1.3),  
638 the time at which a pulse reaches a threshold is dependent upon the size of the pulse.  
639 This is known as the ‘time-walk’ effect and also needs to be accounted for in each PMT.  
640 To calibrate the timing response, a pulse of light with width 0.2ns is emitted into the

641 detector through a diffuser. Two-dimensional distributions of time and pulse height  
642 (or charge) are made for each PMT and are used to calibrate the timing response. This  
643 is performed in-situ during data taking with the light source pulsing at 0.03Hz.

644 The top/bottom water quality asymmetry is measured using the NiCf calibration  
645 data and cross-referencing these results to the “standard PMTs”. The water attenuation  
646 length is continuously measured by the rate of vertically-downgoing cosmic-ray  
647 muons which enter via the top of the tank.

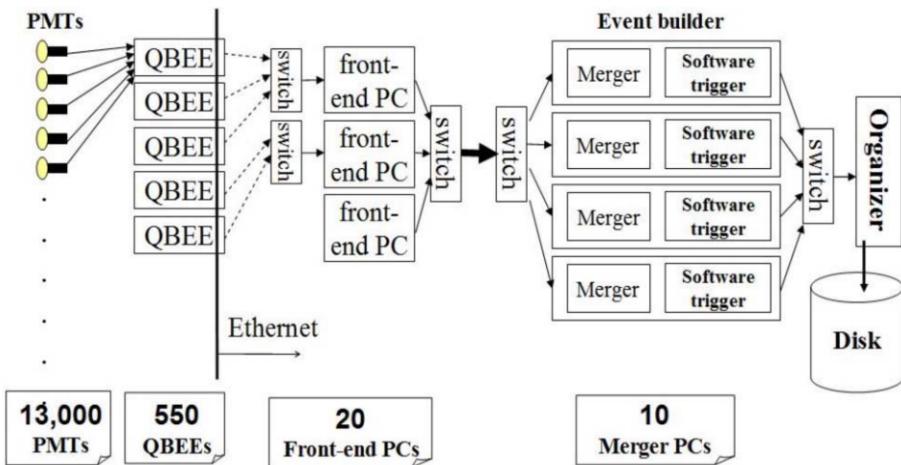
648 Dark noise is the phenomenon where a PMT registers a pulse that is consistent  
649 with a single photoelectron emitted from photon detection despite the PMT being in  
650 complete darkness. This is predominately caused by two processes. Firstly there is  
651 intrinsic dark noise which is where photoelectrons gain enough thermal energy to be  
652 emitted from the photocathode, and secondly, the radioactive decay of contaminants  
653 inside the structure of the PMT. Typical dark noise rate for PMTs used within SK are  
654  $O(3)$ kHz [87]. This is lower than the expected number of photons generated for a ‘high  
655 energy event’ (As described in subsection 3.1.4) but instability in this value can cause  
656 biases in reconstruction. Dark noise is related to the gain of a PMT and is calibrated  
657 using hits inside a time window recorded before an event trigger [98].

### 658 3.1.3. Data Acquisition and Triggering

659 The analysis presented in this thesis only uses the SK-IV period of the SK experiment  
660 so this subsection focuses on the relevant points of the data acquisition and triggering  
661 systems to that SK period. The earlier data acquisition and triggering systems are  
662 documented in [99, 100].

663 Before the SK-IV period started, the existing front-end electronics were replaced  
664 with “QTC-Based Electrons with Ethernet, QBEE” systems [101]. When the QBEE

observes a signal above a 1/4 photoelectron threshold, the charge-to-time (QTC) converter generates a rectangular pulse. The start of the rectangular pulse indicates the time at which the analog photoelectron signal was received and the width of the pulse indicates the total charge integrated throughout the signal. This is then digitized by time-to-digital converters and sent to the “front-end” PCs. The digitized signal from every QBEE is then chronologically ordered and sent to the “merger” PCs. It is the merger PCs that apply the software trigger. Any triggered events are passed to the “organizer” PC. This sorts the data stream of multiple merger PCs into chronologically ordered events which are then saved to disk. The schematic of data flow from PMTs to disk is illustrated in Figure 3.3.



**Figure 3.3.:** Schematic view of the data flow through the data acquisition and online system. Taken from [102].

The software trigger (described in [103]) operates by determining the number of PMT hits within a 200ns sliding window,  $N_{200}$ . This window coincides with the maximum time that a Cherenkov photon would take to traverse the length of the SK tank [100]. For lower energy events that generate fewer photons, this technique is useful for eliminating background processes like dark noise and radioactive decay which would be expected to separate in time. When the value of  $N_{200}$  exceeds some threshold, a software trigger is issued. There are several trigger thresholds used within

the SK-IV period which are detailed in Table 3.2. If one of these thresholds is met, the PMT hits within an extended time window are also read out and saved to disk. In the special case of an event that exceeds the SHE trigger but does not exceed the OD trigger, the AFT trigger looks for delayed coincidences of 2.2MeV gamma rays emitted from neutron capture in a  $535\mu\text{s}$  window after the SHE trigger. A similar but more complex “Wideband Intelligent Trigger (WIT)” has been deployed and is described in [104].

Trigger	Acronym	Condition	Extended time window ( $\mu\text{s}$ )
Super Low Energy	SLE	>34/31 hits	1.3
Low Energy	LE	>47 hits	40
High Energy	HE	>50 hits	40
Super High Energy	SHE	>70/58 hits	40
Outer Detector	OD	>22 hits in OD	N/A

**Table 3.2.:** The trigger thresholds and extended time windows saved around an event which were utilised throughout the SK-IV period. The exact thresholds can change and the values listed here represent the thresholds at the start and end of the SK-IV period.

### 3.1.4. Cherenkov Radiation

Cherenkov light is emitted from any highly energetic charged particle traveling with relativistic velocity,  $\beta$ , greater than the local speed of light in a medium [105]. Cherenkov light is formed at the surface of a cone with characteristic pitch angle,

$$\cos(\theta) = \frac{1}{\beta n}. \quad (3.2)$$

693 where  $n$  is the refractive index of the medium. Consequently, the Cherenkov  
 694 momentum threshold,  $P_{thres}$ , is dependent upon the mass,  $m$ , of the charged particle  
 695 moving through the medium,

$$P_{thres} = \frac{m}{\sqrt{n^2 - 1}} \quad (3.3)$$

696 For water, where  $n = 1.33$ , the Cherenkov threshold momentum and energy for  
 697 various particles are given in Table 3.3. In contrast,  $\gamma$ -rays are detected indirectly via  
 698 the combination of photons generated by Compton scattering and pair production.  
 699 The threshold for detection in the SK detector is typically higher than the threshold  
 700 for photon production. This is due to the fact that the attenuation of photons in the  
 701 water means that typically  $\sim 75\%$  of Cherenkov photons reach the ID PMTs. Then the  
 702 collection and quantum efficiencies described in subsection 3.1.1 result in the number  
 703 of detected photons being lower than the number of photons which reach the PMTs.

Particle	Threshold Momentum (MeV)	Threshold Energy (MeV)
Electron	0.5828	0.7751
Muon	120.5	160.3
Pion	159.2	211.7
Proton	1070.0	1423.1

**Table 3.3.:** The threshold momentum and energy for a particle to generate Cherenkov light in ultrapure water, as calculated in Equation 3.2 in ultrapure water which has refractive index  $n = 1.33$ .

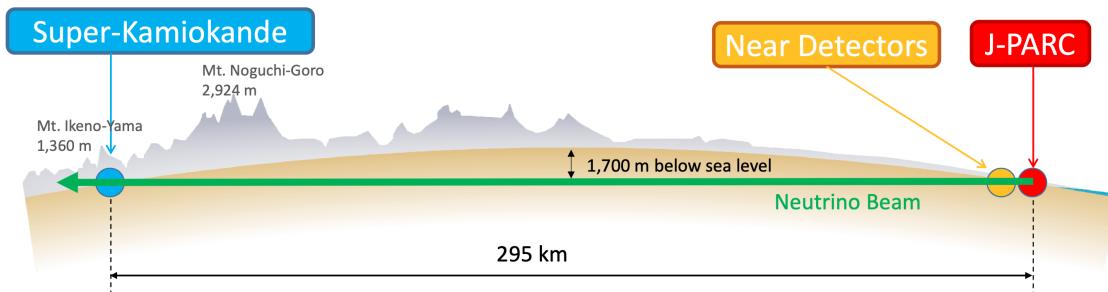
704 The Frank-Tamm equation [106] describes the relationship between the number of  
 705 Cherenkov photons generated per unit length,  $dN/dx$ , the wavelength of the photons  
 706 generated,  $\lambda$ , and the relativistic velocity of the charged particle,

$$\frac{d^2N}{dxd\lambda} = 2\pi\alpha \left(1 - \frac{1}{n^2\beta^2}\right) \frac{1}{\lambda^2}. \quad (3.4)$$

where  $\alpha$  is the fine structure constant. For a 100MeV momentum electron, approximately 330 photons will be produced per centimeter in the  $300\text{nm} \leq \lambda \leq 700\text{nm}$  region which the ID PMTs are most sensitive to [87].

## 3.2. The Tokai to Kamioka Experiment

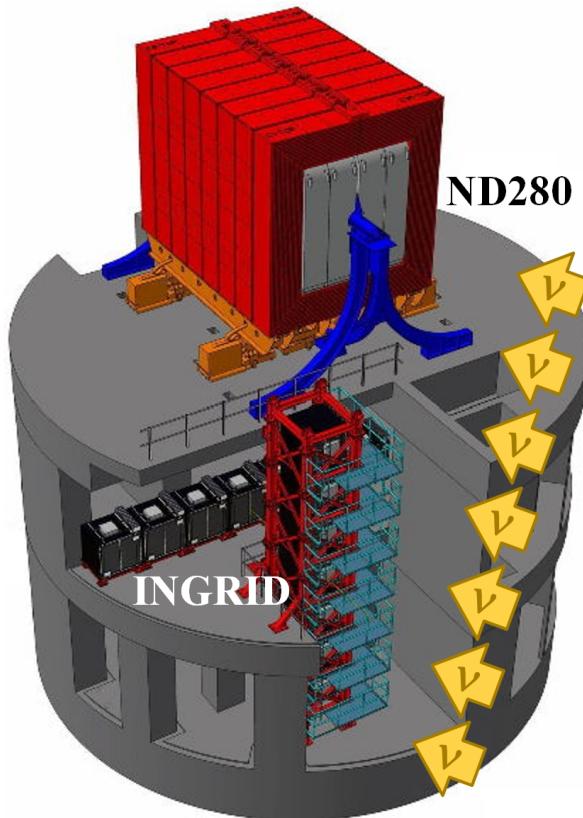
The Tokai to Kamioka (T2K) experiment is a long-baseline neutrino oscillation experiment located in Japan. Proposed in the early 2000s [107, 108] to replace K2K [109], T2K was designed to observe electron neutrino appearance whilst precisely measuring the oscillation parameters associated with muon neutrino disappearance [110]. The experiment consists of a neutrino beam generated at the Japan Proton Accelerator Research Complex (J-PARC), a suite of near detectors situated 280m from the beam target, and the Super Kamiokande far detector positioned at a 295km baseline. The cross-section view of the T2K experiment is drawn in Figure 3.4.



**Figure 3.4.:** The cross-section view of the Tokai to Kamioka experiment illustrating the beam generation facility at J-PARC, the near detector situated at a baseline of 280m and the Super Kamiokande far detector situated 295km from the beam target.

The T2K collaboration makes world-leading measurements of the  $\sin^2(\theta_{23})$ ,  $\Delta m_{23}^2$ , and  $\delta_{CP}$  oscillation parameters. Improvements in the precision and accuracy of parameter estimates are still being made by including new data samples and developing the models which describe the neutrino interactions and detector responses [111]. Electron neutrino appearance was first observed at T2K in 2014 [112] with  $7.3\sigma$  significance.

The near detectors provide constraints on the beam flux and cross-section model parameters used within the oscillation analysis by observing the unoscillated neutrino beam. There are a host of detectors situated in the near detector hall (As illustrated in Figure 3.5): ND280 (subsection 3.2.2), INGRID (subsection 3.2.3), NINJA [113], WAGASCI [114], and Baby-MIND [115]. The latter three are not currently used within the oscillation analysis presented within this thesis.



**Figure 3.5.:** The near detector suite for the T2K experiment showing the ND280 and INGRID detectors. The distance between the detectors and the beam target is 280m.

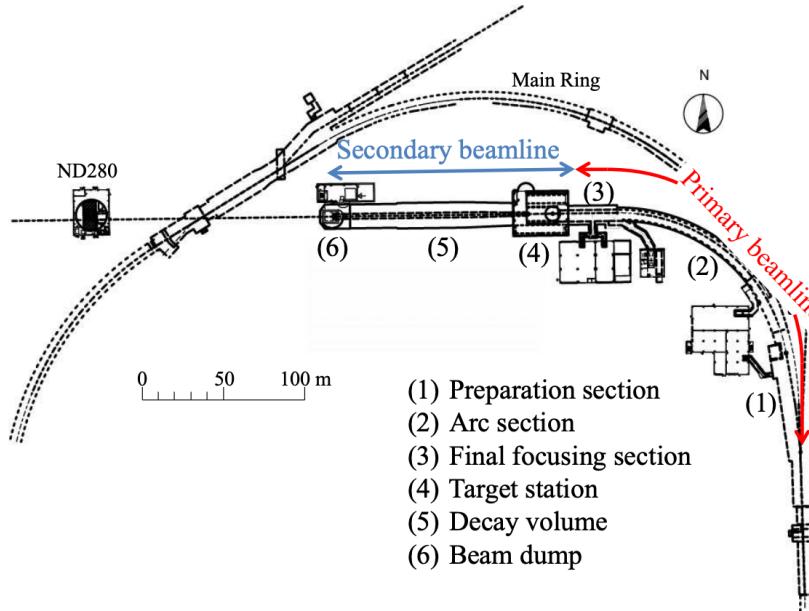
Whilst this thesis presents the ND280 in terms of its purpose for the oscillation analysis, the detector can also make many cross-section measurements at neutrino energies of  $O(1)$ GeV for the different targets within the detector [116, 117]. These measurements are of equal importance as they can lead the way in determining the model parameters used in the interaction models for the future high-precision era of neutrino physics.

DB: Discuss BANFF, PTheta, MaCh3 and covariance

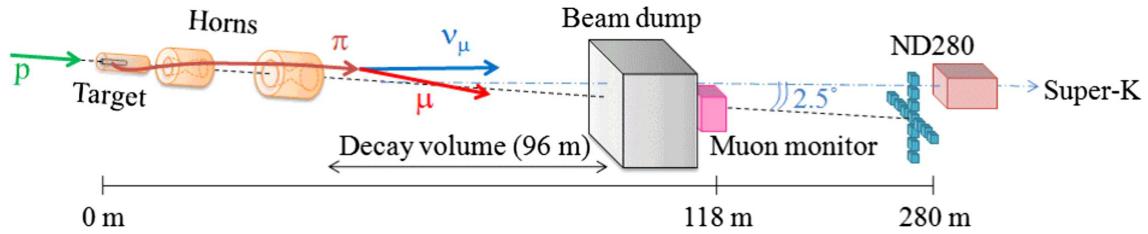
### 3.2.1. The Neutrino Beam

The neutrino beam used within the T2K experiment is described in [59, 118] and summarised below. The accelerating facility at J-PARC is composed of two sections; the primary and secondary beamlines. Figure 3.6 illustrates a schematic of the beamline, focusing mostly on the components of the secondary beamline. The primary beamline has three accelerators that progressively accelerate protons; a linear accelerator, a rapid-cycling synchrotron, and the main-ring (MR) synchrotron. Once fully accelerated by the MR, the protons have a kinetic energy of 30GeV. Eight bunches of these protons, separated by 500ns, are extracted per “spill” from the MR and directed towards a graphite target (a rod of length 91.4cm and diameter 2.6cm). Spills are extracted at 0.5Hz with  $\sim 3 \times 10^{14}$  protons contained per spill.

The secondary beamline consists of three main components: the target station, the decay volume, and the beam dump. The target station is comprised of the target, beam monitors, and three magnetic focusing horns. The proton beam interacts with the graphite target to form a secondary beam of mostly pions and kaons. The secondary beam travels through a 96m long decay volume, generating neutrinos through the following decays [59],



(a) Primary and secondary beamline



(b) Secondary beamline

**Figure 3.6.:** Top panel: Bird's eye view of the most relevant part of primary and secondary beamline used within the T2K experiment. The primary beamline is the main-ring proton synchrotron, kicker magnet, and graphite target. The secondary beamline consists of the three focusing horns, decay volume, and beam dump. Figure taken from [118]. Bottom panel: The side-view of the secondary beamline including the focusing horns, beam dump and neutrino detectors. Figure taken from [119].

$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$

$$\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$$

$$K^+ \rightarrow \mu^+ + \nu_\mu$$

$$K^- \rightarrow \mu^- + \bar{\nu}_\mu$$

$$\rightarrow \pi^0 + e^+ + \nu_e$$

$$\rightarrow \pi^0 + e^- + \bar{\nu}_e$$

$$\rightarrow \pi^0 + \mu^+ + \nu_\mu$$

$$\rightarrow \pi^0 + \mu^- + \bar{\nu}_\mu$$

$$K_L^0 \rightarrow \pi^- + e^+ + \nu_e$$

$$K_L^0 \rightarrow \pi^+ + e^- + \bar{\nu}_e$$

$$\rightarrow \pi^- + \mu^+ + \nu_\mu$$

$$\rightarrow \pi^+ + \mu^- + \bar{\nu}_\mu$$

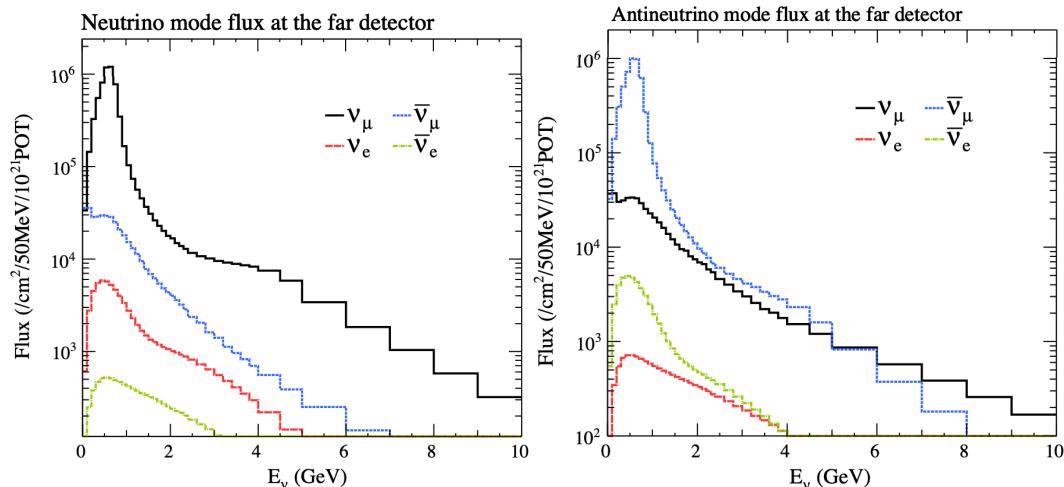
$$\mu^+ \rightarrow e^+ + \bar{\nu}_\mu + \nu_e$$

$$\mu^- \rightarrow e^- + \nu_\mu + \bar{\nu}_e$$

The electrically charged component of the secondary beam is focused towards the far detector by the three magnetic horns. These horns direct charged particles of a particular polarity towards SK whilst defocusing the oppositely charged particles. This allows a mostly neutrino or mostly antineutrino beam to be used within the experiment, denoted as “forward horn current (FHC)” or “reverse horn current (RHC)” respectively.

Figure 3.7 illustrates the different contributions to the FHC and RHC neutrino flux.

The low energy flux is dominated by the decay of pions whereas kaon decay becomes the dominant source of neutrinos for  $E_\nu > 3\text{GeV}$ . The “wrong-sign” component, which is the  $\bar{\nu}_\mu$  background in a  $\nu_\mu$  beam, and the intrinsic irreducible  $\nu_e$  background, are predominantly due to muon decay for  $E_\nu < 2\text{GeV}$ . As the antineutrino production cross-section is smaller than the neutrino cross-section, the wrong-sign component is more dominant in the RHC beam as compared to that in the FHC beam.



**Figure 3.7.:** The Monte Carlo prediction of the energy spectrum for each flavour of neutrino ( $\nu_e$ ,  $\bar{\nu}_e$ ,  $\nu_\mu$  and  $\bar{\nu}_\mu$ ) in the neutrino dominated beam FHC mode (Left) and antineutrino dominated beam RHC mode (Right) expected at SK. Taken from [120].

The beam dump, situated at the end of the decay volume, stops all charged particles other than highly energetic muons ( $p_\mu > 5\text{GeV}$ ). The MuMon detector monitors the

<sup>770</sup> penetrating muons to determine the beam direction and intensity which is used to  
<sup>771</sup> constrain some of the beam flux systematics within the analysis [119, 121].

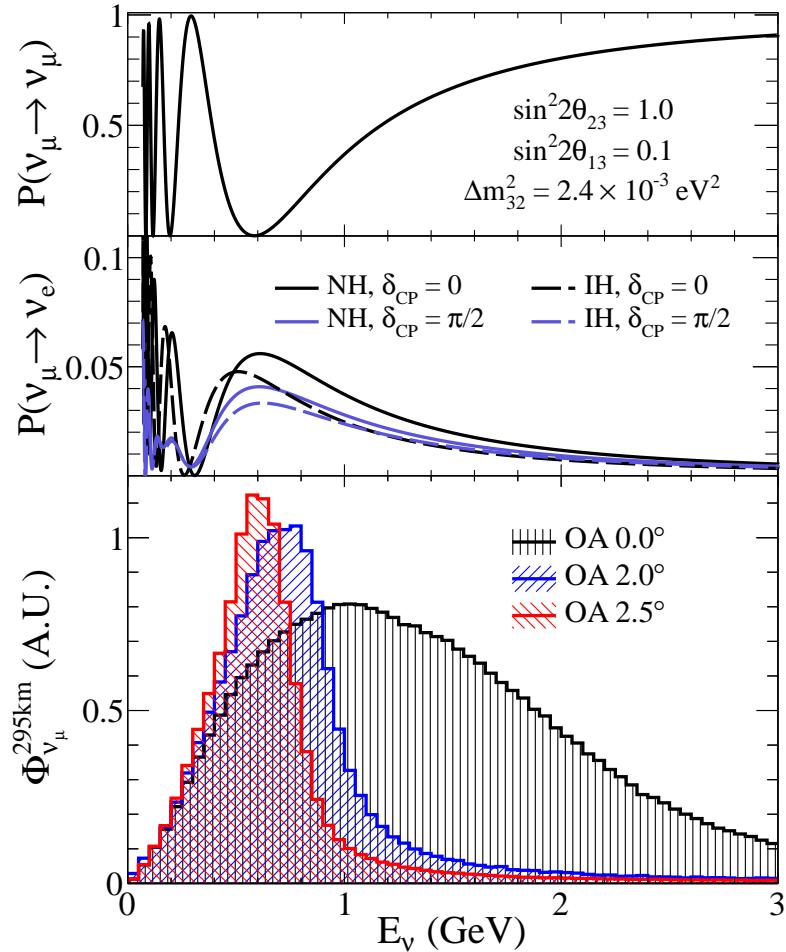
<sup>772</sup> The T2K experiment uses an off-axis beam to narrow the neutrino energy distribution.  
<sup>773</sup> This was the first implementation of this technique in a long-baseline neutrino  
<sup>774</sup> oscillation experiment after its original proposal [122]. Pion decay,  $\pi \rightarrow \mu + \nu_\mu$ , is a  
<sup>775</sup> two-body decay. Consequently, the neutrino energy,  $E_\nu$ , can be determined based on  
<sup>776</sup> the pion energy,  $E_\pi$ , and the angle at which the neutrino is emitted,  $\theta$ ,

$$E_\nu = \frac{m_\pi^2 - m_\mu^2}{2(E_\pi - p_\pi \cos(\theta))}, \quad (3.5)$$

<sup>777</sup> where  $m_\pi$  and  $m_\mu$  are the mass of the pion and muon respectively. For a fixed  
<sup>778</sup> energy pion, the neutrino energy distribution is dependent upon the angle at which the  
<sup>779</sup> neutrinos are observed from the initial pion beam direction. For the 295km baseline at  
<sup>780</sup> T2K,  $E_\nu = 0.6\text{GeV}$  maximises the electron neutrino appearance probability,  $P(\nu_\mu \rightarrow \nu_e)$ ,  
<sup>781</sup> whilst minimising the muon disappearance probability,  $P(\nu_\mu \rightarrow \nu_\mu)$ . Figure 3.8  
<sup>782</sup> illustrates the neutrino energy distribution for a range of off-axis angles, as well as the  
<sup>783</sup> oscillation probabilities most relevant to T2K.

### <sup>784</sup> 3.2.2. The Near Detector at 280m

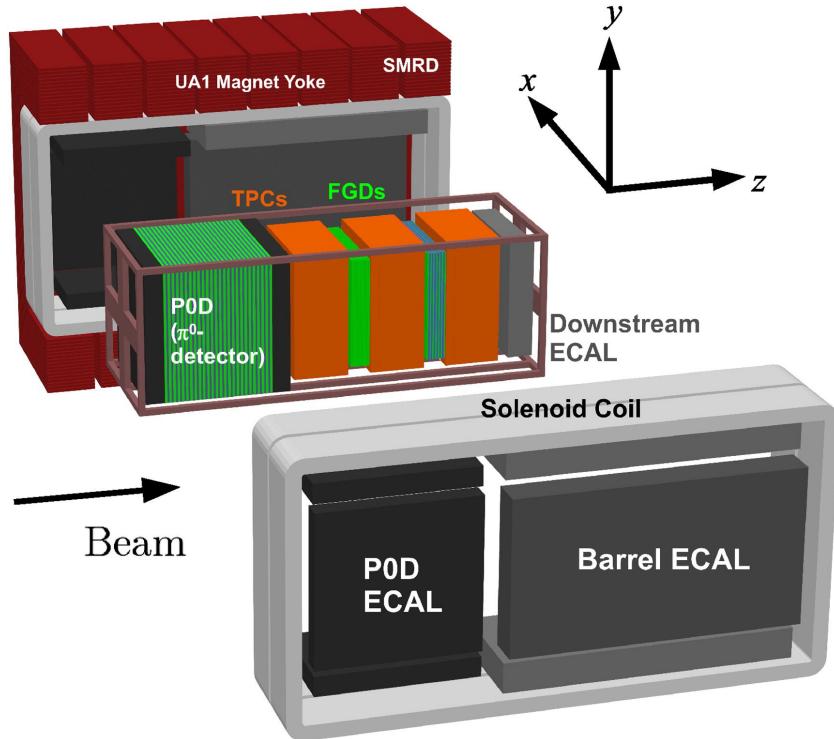
<sup>785</sup> Whilst all the near detectors are situated in the same “pit” located at 280m from the  
<sup>786</sup> beamline, the “ND280” detector is the off-axis detector which is situated at the same  
<sup>787</sup> off-axis angle as the Super-Kamiokande far detector. It has two primary functions;  
<sup>788</sup> firstly it measures the neutrino flux and secondly it counts the event rates of different  
<sup>789</sup> types of neutrino interactions. Both of these constrain the flux and cross-section



**Figure 3.8.:** Top panel: T2K muon neutrino disappearance probability as a function of neutrino energy. Middle panel: T2K electron neutrino appearance probability as a function of neutrino energy. Bottom panel: The neutrino flux distribution for three different off-axis angles (Arbitrary units) as a function of neutrino energy.

790 systematics invoked within the model for a more accurate prediction of the expected  
791 event rate at the far detector.

792 As illustrated in Figure 3.9, the ND280 detector consists of several sub-detectors.  
793 The most important part of the detector for this analysis is the tracker region. This is  
794 comprised of two time projection chambers (TPCs) sandwiched between three fine  
795 grain detectors (FGDs). The FGDs contain both hydrocarbon plastics and water tar-  
796 gets for neutrino interactions and provide track reconstruction near the interaction  
797 vertex. The emitted charged particles can then propagate into the TPCs which pro-  
798 vide particle identification and momentum reconstruction. The FGDs and TPCs are



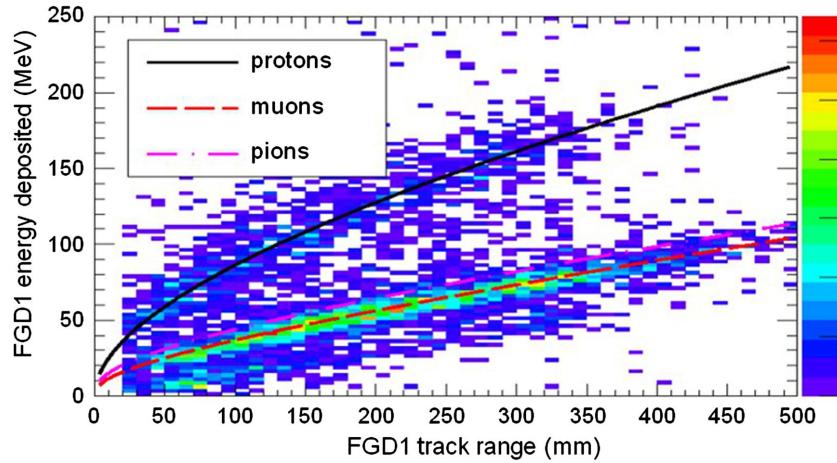
**Figure 3.9.:** The components of the ND280 detector. The neutrino beam travels from left to right. Taken from [118].

799 further described in subsubsection 3.2.2.1 and subsubsection 3.2.2.2 respectively. The  
 800 electromagnetic calorimeter (ECAL) encapsulates the tracker region alongside the  $\pi^0$   
 801 detector (P0D). The ECAL measures the deposited energy from photons emitted from  
 802 interactions within the FGD. The P0D constrains the cross-section of neutral current  
 803 interactions which generate neutral pions, which is one of the largest backgrounds in  
 804 the electron neutrino appearance oscillation channel. The P0D and ECAL detectors  
 805 are detailed in subsubsection 3.2.2.3 and subsubsection 3.2.2.4 respectively. The entire  
 806 detector is located within a large yoke magnet which produces a 0.2T magnetic field.  
 807 This design of the magnet also includes a scintillating detector called the side muon  
 808 range detector (SMRD) which is used to track high-angle muons as well as acting as a  
 809 cosmic veto. The SMRD is described in subsubsection 3.2.2.5.

810    **3.2.2.1. Fine Grained Detectors**

811    The T2K tracker region is comprised of two fine grained detectors (FGD) and three  
812    Time Projection Chambers (TPC). A detailed description of the FGD design, construc-  
813    tion, and assembly is found in [123] and summarised below. The FGDs are the primary  
814    target for neutrino interactions with a mass of 1.1 tonnes per FGD. Alongside this,  
815    the FGDs are designed to be able to track short-range particles which do not exit the  
816    FGD. Typically, short-range particles are low momentum and are observed as tracks  
817    that deposit a large amount of energy per unit length. This means the FGD needs  
818    good granularity to resolve these particles. The FGDs have the best timing resolution  
819    ( $\sim 3\text{ns}$ ) of any of the sub-detectors of the ND280 detector. As such, the FGDs are  
820    used for time of flight measurements to distinguish forward going positively charged  
821    particles from backward going negatively charged particles. Finally, any tracks which  
822    pass through multiple sub-detectors are required to be track matched to the FGD.

823       Both FGDs are made from square scintillator planes of side length 186cm and  
824    width 2.02cm. Each plane consists of two layers of 192 scintillator bars in an X or Y  
825    orientation. A wavelength shifting fiber is threaded through the center of each bar and  
826    is read out by a multi-pixel photon counter (MPPC). FGD1 is the most upstream of  
827    the two FGDs and contains 15 planes of carbon plastic scintillator which is a common  
828    target in external neutrino scattering data. As the far detector is a pure water target, 7  
829    of the 15 scintillator planes in FGD2 have been replaced with a hybrid water-scintillator  
830    target. Due to the complexity of the nucleus, nuclear effects can not be extrapolated  
831    between different nuclei. Therefore having the ability to take data on one target which  
832    is the same as external data and another target which is the same as the far detector  
833    target is beneficial for reliable model parameter estimates.



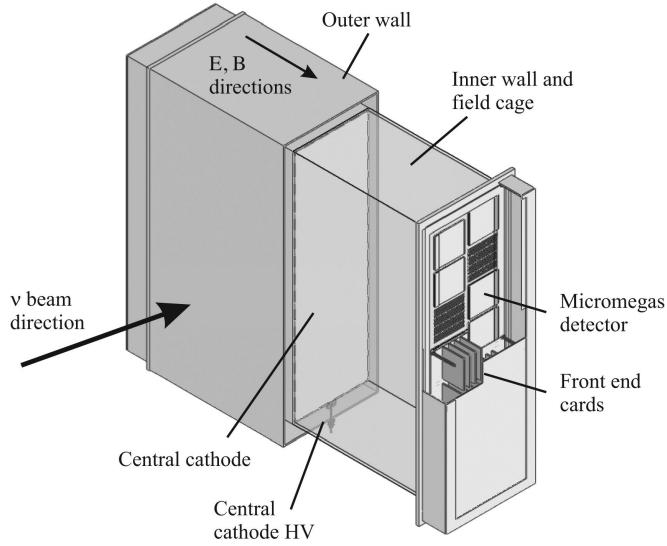
**Figure 3.10.:** Comparison of data to Monte Carlo prediction of integrated deposited energy as a function of track length for particles that stopped in FGD1. Taken from [123].

The integrated deposited energy is used for particle identification. The FGD can distinguish protons from other charged particles by comparing the integrated deposited energy from data to Monte Carlo prediction as seen in Figure 3.10.

### 3.2.2.2. Time Projection Chambers

The majority of particle identification and momentum measurements within ND280 are provided by three Time Projection Chambers (TPCs) [124]. The TPCs are located on either side of the FGDs. They are located inside of the magnetic field meaning the momentum of a charged particle can be determined from the bending of the track.

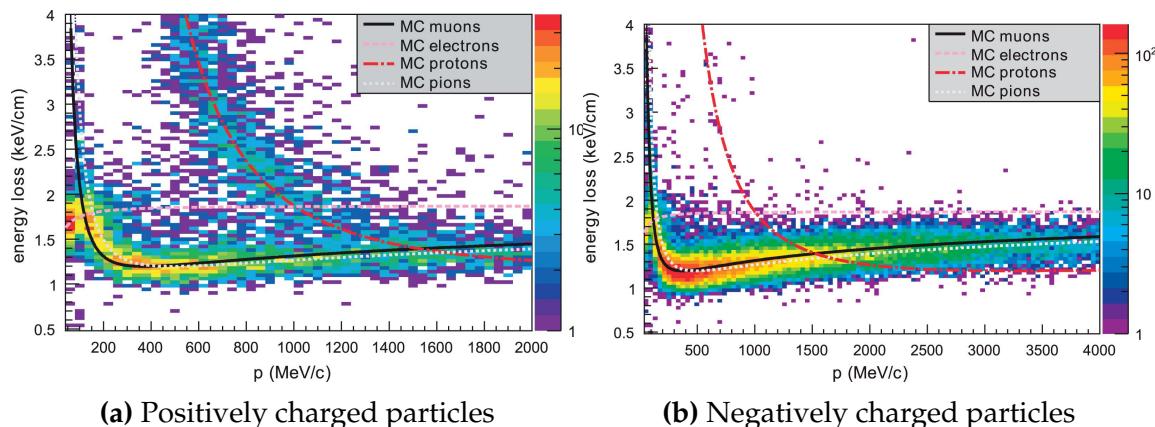
Each TPC module consists of two gas-tight boxes, as shown in Figure 3.11, which are made of non-magnetic material. The outer box is filled with CO<sub>2</sub> which acts as an electrical insulator between the inner box and the ground. The inner box forms the field cage which produces a uniform electric drift field of  $\sim 275\text{V}/\text{cm}$  and is filled with an argon gas mixture. Charged particles moving through this gas mixture ionize the gas and the ionised charge is drifted towards micromegas detectors which measure the ionization charge. The time and position information in the readout allows a three-dimensional image of the neutrino interaction.



**Figure 3.11.:** Schematic design of a Time Projection Chamber detector. Taken from [124].

The particle identification of tracks that pass through the TPCs is performed using

dE/dx measurements. Figure 3.12 illustrates the data to Monte Carlo distributions of the energy lost by a charged particle passing through the TPC as a function of the reconstructed particle momentum. The resolution is  $7.8 \pm 0.2\%$  meaning that electrons and muons can be distinguished. This allows reliable measurements of the intrinsic  $\nu_e$  component of the beam.



**Figure 3.12.:** The distribution of energy loss as a function of reconstructed momentum for charged particles passing through the TPC, comparing data to Monte Carlo prediction. Taken from [124].

**<sup>856</sup> 3.2.2.3.  $\pi^0$  Detector**

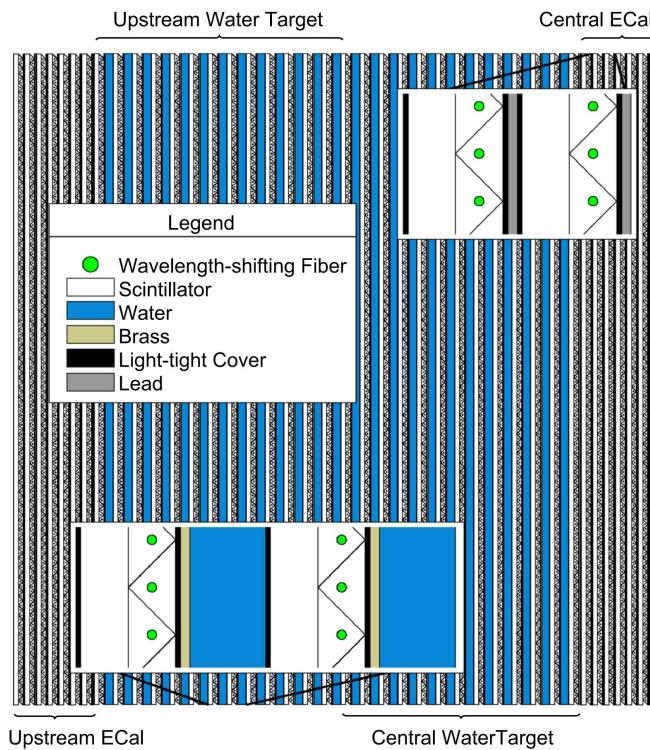
<sup>857</sup> If one of the  $\gamma$ -rays from a  $\pi^0 \rightarrow 2\gamma$  decay is missed at the far detector, the recon-  
<sup>858</sup> struction will determine that event to be a charge current  $\nu_e$ -like event. This is one of  
<sup>859</sup> the main backgrounds hindering the electron neutrino appearance searches. The  $\pi^0$   
<sup>860</sup> detector (P0D) measures the cross-section of the neutral current induced neutral pion  
<sup>861</sup> production on a water target to constrain this background.

<sup>862</sup> The P0D is a cube of approximately 2.5m length consisting of layers of scintillating  
<sup>863</sup> bars, brass and lead sheets, and water bags as illustrated in Figure 3.13. Two electro-  
<sup>864</sup> magnetic calorimeters are positioned at the most upstream and most downstream  
<sup>865</sup> position in the sub-detector and the water target is situated in between them. The  
<sup>866</sup> scintillator layers are built from two triangular bars orientated in opposite directions  
<sup>867</sup> to form a rectangular layer. Each triangular scintillator bar is threaded with optical  
<sup>868</sup> fiber which is read out by MPPCs. The high-Z brass and lead regions produce electron  
<sup>869</sup> showers from the photons emitted in  $\pi^0$  decay.

<sup>870</sup> The sub-detector can generate measurements of NC1 $\pi^0$  cross-sections on a water  
<sup>871</sup> target by measuring the event rate both with and without the water target, with the  
<sup>872</sup> cross-section on a water target being determined as the difference. The total active  
<sup>873</sup> mass is 16.1 tonnes when filled with water and 13.3 tonnes when empty.

**<sup>874</sup> 3.2.2.4. Electromagnetic Calorimeter**

<sup>875</sup> The electromagnetic calorimeter [126] (ECal) encapsulates the P0D and tracking sub-  
<sup>876</sup> detectors. Its primary purpose is to aid  $\pi^0$  reconstruction from any interaction in  
<sup>877</sup> the tracker. To do this, it measures the energy and direction of photon showers from  
<sup>878</sup>  $\pi^0 \rightarrow 2\gamma$  decay. It can also distinguish pion and muon tracks depending on the shape  
<sup>879</sup> of the photon shower deposited.



**Figure 3.13.:** A schematic of the P0D side-view. Taken from [125].

The ECal is comprised of three sections; the P0D ECal which surrounds the P0D, the barrel ECal which encompasses the tracking region, and the downstream ECal which is situated downstream of the tracker region. The barrel and downstream ECals are tracking calorimeters that focus on electromagnetic showers from high-angle particles emitted from the tracking sub-detectors. Particularly in the TPC, high-angle tracks (those which travel perpendicularly to the beam-axis) can travel along a single scintillator bar resulting in very few hits. The width of the barrel and downstream ECal corresponds to  $\sim 11$  electron radiation lengths to ensure a significant amount of the  $\pi^0$  energy is contained. As the P0D has its own calorimetry which reconstructs showers, the P0D ECal determines the energy which escapes the P0D.

Each ECal is constructed of multiple layers of scintillating bars sandwiched between lead sheets. The scintillating bars are threaded with optical fiber and read out by MPPCs. Each sequential layer of the scintillator is orientated perpendicular to the previous which allows a three dimensional event reconstruction. The target mass

894 of the P0D ECal, barrel ECal, and downstream ECal are 1.50, 4.80 and 6.62 tonnes  
895 respectively.

896 **3.2.2.5. Side Muon Range Detector**

897 As illustrated in Figure 3.9, the ECal, FGDs, P0D, and TPCs are enclosed within the  
898 UA1 magnet. Originally designed for the NOMAD [127] experiment and reconditioned  
899 for use in the T2K experiment [128], the UA1 magnet provides a uniform horizontal  
900 magnetic field of 0.2T with an uncertainty of  $2 \times 10^{-4}$ T.

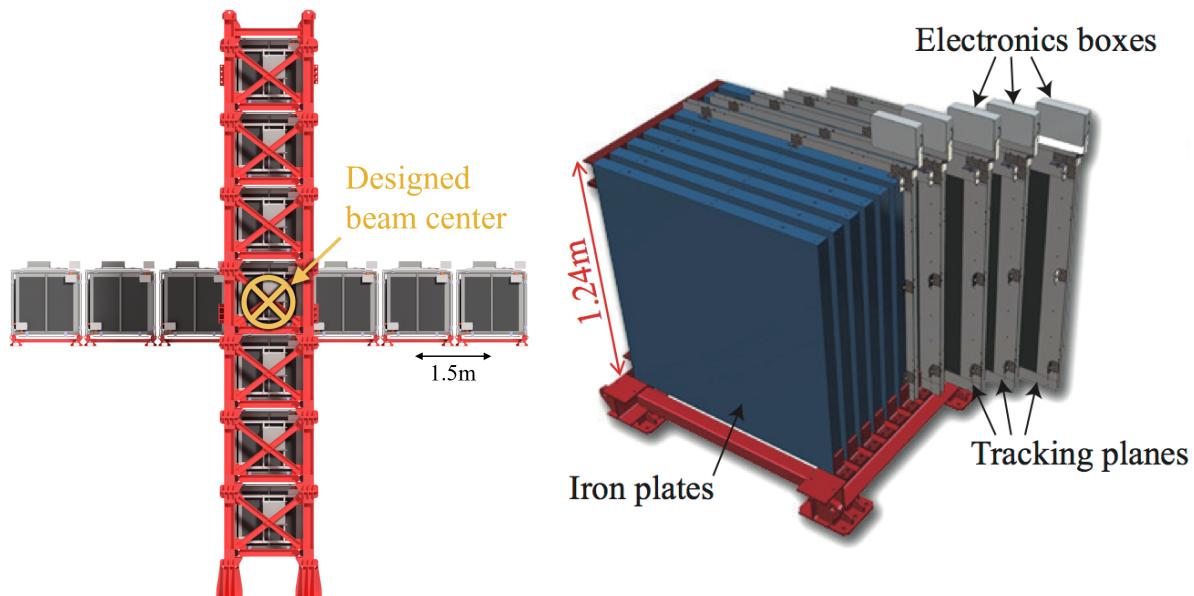
901 Built into the UA1 magnet, the side muon range detector (SMRD) [129] monitors  
902 high-energy muons which leave the tracking region and permeate through the ECal.  
903 It additionally acts as a cosmic muon veto and trigger.

904 **3.2.3. The Interactive Neutrino GRID**

905 The Interactive Neutrino GRID (INGRID) detector is situated within the same “pit” as  
906 the other near detectors. It is aligned with the beam in the “on-axis” position and mea-  
907 sures the beam direction, spread, and intensity. The detector was originally designed  
908 with 16 identical modules [118] (two modules have since been decommissioned) and a  
909 “proton” module. The design of the detector is cross-shaped with length and height  
910 10m × 10m as illustrated in Figure 3.14.

911 Each module is composed of iron sheets interlaced with eleven tracking scintillator  
912 planes for a total target mass of 7.1 tonnes per module. The scintillator design is an X-Y  
913 pattern of 24 bars in both orientations, where each bar contains wave-length shifting  
914 fibers which are connected to multi-pixel photon counters (MPPCs). Each module is  
915 encapsulated inside veto planes to aid the rejection of charged particles entering the  
916 module.

917     The proton module is different from the other modules in that it consists of entirely  
 918   scintillator planes with no iron target. The scintillator bars are also smaller than those  
 919   used in the other modules to increase the granularity of the detector and improve  
 920   tracking capabilities. The module sits in the center of the beamline and is designed to  
 921   give precise measurements of quasi-elastic charged current interactions to evaluate  
 922   the performance of the Monte Carlo simulation of the beamline.



**Figure 3.14.:** Left panel: The Interactive Neutrino GRID on-axis Detector. 14 modules are arranged in a cross-shape configuration, with the center modules being directly aligned with the on-axis beam. Right panel: The layout of a single module of the INGRID detector. Both figures are recreated from [118].

923     The INGRID detector can measure the beam direction to an uncertainty of 0.4mrad  
 924   and the beam center within a resolution of 10cm [118]. The beam direction in both the  
 925   vertical and horizontal directions is discussed in [130] and it is found to be in good  
 926   agreement with the MUMON monitor described in subsection 3.2.1.

<sub>927</sub> **Chapter 4**

<sub>928</sub> **Bayesian Statistics and Markov Chain  
Monte Carlo Techniques**

<sub>930</sub> This thesis presents a Bayesian oscillation analysis. To extract the oscillation parameters,  
<sub>931</sub> a Markov Chain Monte Carlo (MCMC) method is used. This chapter explains  
<sub>932</sub> the theory of how parameter estimates can be determined using this technique and  
<sub>933</sub> condenses the material found in the literature [131–134].

<sub>934</sub> The oscillation parameter determination presented within this thesis is built upon  
<sub>935</sub> a simultaneous fit to neutrino beam data in the near detector, beam data at SK and  
<sub>936</sub> atmospheric data at SK. In total, there are four oscillation parameters of interest  
<sub>937</sub> ( $\sin^2(\theta_{23})$ ,  $\sin^2(\theta_{13})$ ,  $\Delta m_{23}^2$ , and  $\delta_{CP}$ ), two oscillation parameters to which this study  
<sub>938</sub> will not be sensitive ( $\sin^2(\theta_{12})$ ,  $\Delta m_{12}^2$ ) and many nuisance parameters that control the  
<sub>939</sub> systematic uncertainty models invoked within this study.

<sub>940</sub> The MCMC technique generates a multi-dimensional probability distribution across  
<sub>941</sub> all of the model parameters used in the fit. To determine the parameter estimate of a  
<sub>942</sub> single parameter, this multi-dimensional object is integrated over all other parameters.  
<sub>943</sub> This process is called Marginalisation and is further described in subsection 4.3.1.  
<sub>944</sub> Monte Carlo techniques approximate the probability distribution of each parameter  
<sub>945</sub> within the limit of generating infinite samples. As ever, generating a large number of  
<sub>946</sub> samples is time and resource-dependent. Therefore, an MCMC technique is utilised  
<sub>947</sub> within this analysis to reduce the required number of steps to sufficiently sample the  
<sub>948</sub> parameter space. This technique is described in further detail in subsection 4.2.1.

## <sup>949</sup> 4.1. Bayesian Statistics

<sup>950</sup> Bayesian inference treats observable data,  $D$ , and model parameters,  $\vec{\theta}$ , on equal  
<sup>951</sup> footing such that a probability model of both data and parameters is required. This is  
<sup>952</sup> the joint probability distribution  $P(D, \vec{\theta})$  and can be described by the prior distribution  
<sup>953</sup> for model parameters  $P(\vec{\theta})$  and the likelihood of the data given the model parameters  
<sup>954</sup>  $P(D|\vec{\theta})$ ,

$$P(D, \vec{\theta}) = P(D|\vec{\theta})P(\vec{\theta}). \quad (4.1)$$

<sup>955</sup> The prior distribution,  $P(\vec{\theta})$ , describes all previous knowledge about the parameters  
<sup>956</sup> within the model. For example, if the risk of developing health problems is known  
<sup>957</sup> to increase with age, the prior distribution would describe the increase. For the  
<sup>958</sup> purpose of this analysis, the prior distribution is typically the best-fit values taken  
<sup>959</sup> from external data measurements with a Gaussian uncertainty. The prior distribution  
<sup>960</sup> can also contain correlations between model parameters. In an analysis using Monte  
<sup>961</sup> Carlo techniques, the likelihood of measuring some data assuming some set of model  
<sup>962</sup> parameters is calculated by comparing the Monte Carlo prediction generated at that  
<sup>963</sup> particular set of model parameters to the data.

<sup>964</sup> It is parameter estimation that is important for this analysis and as such, we apply  
<sup>965</sup> Bayes' theorem [135] to calculate the probability for each parameter to have a certain  
<sup>966</sup> value given the observed data,  $P(\vec{\theta}|D)$ , which is known as the posterior distribution  
<sup>967</sup> (often termed the posterior). This can be expressed as

$$P(\vec{\theta}|D) = \frac{P(D|\vec{\theta})P(\vec{\theta})}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}}. \quad (4.2)$$

968 The denominator in Equation 4.2 is the integral of the joint probability distribution

969 over all values of all parameters used within the fit. For brevity, we say that the  
 970 posterior distribution is

$$P(\vec{\theta}|D) \propto P(D|\vec{\theta})P(\vec{\theta}). \quad (4.3)$$

971 In subsection 4.3.1, we see that for the cases used within this analysis, it is reason-

972 able to know the posterior to some normalisation constant.

## 973 4.2. Monte Carlo Simulation

974 Monte Carlo techniques are used to numerically solve a complex problem that does  
 975 not necessarily have an analytical solution. These techniques rely on building a large  
 976 ensemble of samples from an unknown distribution and then using the ensemble to  
 977 approximate the properties of the distribution.

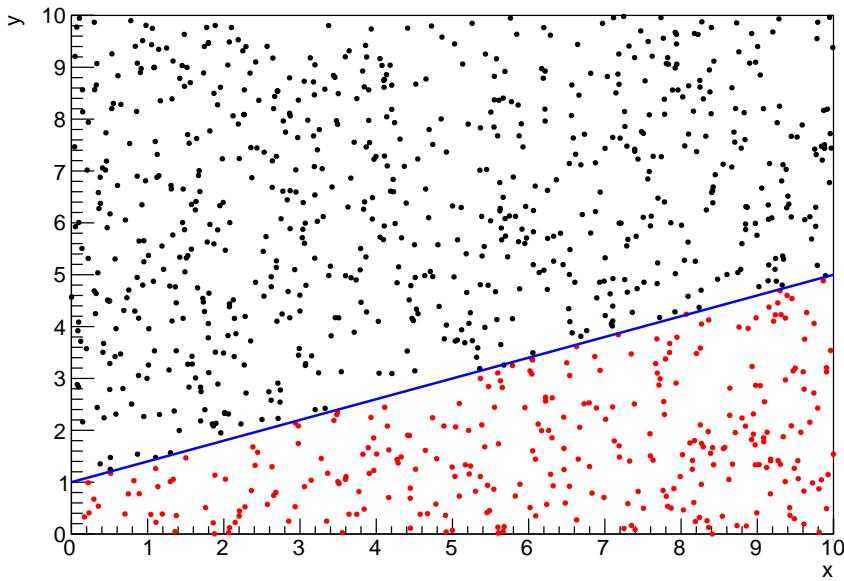
978 An example that uses Monte Carlo techniques is to calculate the area underneath  
 979 a curve. For example, take the problem of calculating the area under a straight line  
 980 with gradient  $M = 0.4$  and intercept  $C = 1.0$ . Analytically, one can calculate the area  
 981 under the line is equal to 30 units for  $0 \leq x \leq 10$ . Using Monte Carlo techniques,  
 982 one can calculate the area under this line by throwing many random values for the  $x$   
 983 and  $y$  components of each sample and then calculating whether that point falls below

the line. The area can then be calculated by the ratio of points below the line to the total number of samples thrown multiplied by the total area in which samples were scattered. The study is shown in Figure 4.1 highlights this technique and finds the area under the curve to be 29.9 compared to an analytical solution of 30.0. The deviation of the numerical to analytical solution can be attributed to the number of samples used in the study. The accuracy of the approximation in which the properties of the Monte Carlo samples replicate those of the desired distribution is dependent on the number of samples used. Replicating this study with a differing number of Monte Carlo samples used in each study (As shown in Figure 4.2) highlights how the Monte Carlo techniques are only accurate within the limit of a high number of samples.

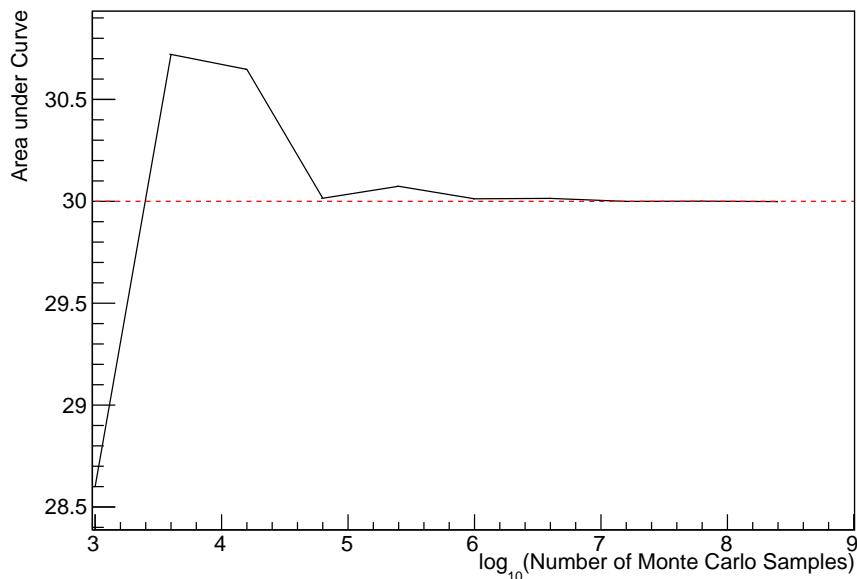
Whilst the above example has an analytical solution, these techniques are just as applicable to complex solutions. Clearly, any numerical solution is only as useful as its efficiency. As discussed, the accuracy of the Monte Carlo technique is dependent upon the number of samples generated to approximate the properties of the distribution. Furthermore, if the positions at which the samples are evaluated are not ‘cleverly’ picked, the efficiency of the Monte Carlo technique significantly drops. Given the example in Figure 4.1, if the region in which the samples are scattered significantly extends passed the region of interest, many calculations will be calculated but do not add to the ability of the Monte Carlo technique to achieve the correct result. For instance, any sample evaluated at a  $y \geq 5$  could be removed without affecting the final result. This does bring in an aspect of the ‘chicken and egg’ problem in that to achieve efficient sampling, one needs to know the distribution beforehand.

#### 4.2.1. Markov Chain Monte Carlo

This analysis utilises a multi-dimensional probability distribution, with some dimensions being significantly more constrained than others. This could be from prior



**Figure 4.1.:** Example of using Monte Carlo techniques to find the area under the blue line. The gradient and intercept of the line are 0.4 and 1.0 respectively. The area found to be under the curve using one thousand samples is 29.9 units.



**Figure 4.2.:** The area under a line of gradient 0.4 and intercept 1.0 for the range  $0 \leq x \leq 10$  as calculated using Monte Carlo techniques as a function of the number of samples used in each repetition. The analytical solution to the area is 30 units as given by the red line.

knowledge of parameter distributions from external data or un-physical regions in which parameters can not exist. Consequently, the Monte Carlo techniques used need to be as efficient as possible. For this analysis, the Markov Chain Monte Carlo (MCMC) technique is chosen. An MCMC technique is a Monte Carlo technique that uses a Markov chain to select which points at which to sample the parameter distribution. This technique performs a semi-random stochastic walk through the allowable parameter space. This builds a posterior distribution which has the property that the density of sampled points is proportional to the probability density of that parameter. This does mean that the samples produced by this technique are not statistically independent but they will cover the space of the distribution.

A Markov chain functions by selecting the position of step  $\vec{x}_{i+1}$  based on the position of  $\vec{x}_i$ . The space in which the Markov chain selects samples is dependent upon the total number of parameters utilised within the fit, where a discrete point in this space is described by the N-dimensional space  $\vec{x}$ . In a perfectly operating Markov chain, the position of the next step depends solely on the previous step and not on the further history of the chain ( $\vec{x}_0, \vec{x}_1$ , etc.). However, in solving the multi-dimensionality of the fit used within this analysis, each step becomes correlated with several of the steps preceding itself. This behaviour is further explained in subsection 4.2.3. Providing the MCMC chain is well optimised, it will begin to converge towards a unique stationary distribution. The period between the chain's initial starting point and the convergence to the unique stationary distribution is colloquially known as the burn-in period. This is discussed further in subsection 4.2.3. Once the chain reaches the stationary distribution, all points sampled after that point will look like samples from that distribution.

1033      Further details of the theories underpinning MCMC techniques are discussed  
1034    in [132] but can be summarised by the requirement that the chain satisfies the three  
1035    ‘regularity conditions’:

- 1036    • Irreducibility: From every position in the parameter space  $\vec{x}$ , there must exist a  
1037       non-zero probability for every other position in the parameter space to be reached.
- 1038    • Recurrence: Once the chain arrives at the stationary distribution, every step fol-  
1039       lowing from that position must be samples from the same stationary distribution.
- 1040    • Aperiodicity: The chain must not repeat the same sequence of steps at any point  
1041       throughout the sampling period.

1042      The output of the chain after burn-in (ie. the sampled points after the chain  
1043    has reached the stationary distribution) can be used to approximate the posterior  
1044    distribution and model parameters  $\vec{\theta}$ . To achieve the requirement that the unique  
1045    stationary distribution found by the chain be the posterior distribution, one can use  
1046    the Metropolis-Hastings algorithm. This guides the stochastic process depending on  
1047    the likelihood of the current proposed step compared to that of the previous step.  
1048    Implementation and other details of this technique are discussed in subsection 4.2.2.

#### 1049 4.2.2. Metropolis-Hastings Algorithm

1050    As a requirement for MCMCs, the Markov chain implemented in this technique must  
1051    have a unique stationary distribution that is equivalent to the posterior distribution.  
1052    To ensure this requirement and that the regularity conditions are met, this analysis  
1053    utilises the Metropolis-Hastings (MH) algorithm [136,137]. For the  $i^{th}$  step in the chain,  
1054    the MH algorithm determines the position in the parameter space to which the chain  
1055    moves to based on the current step,  $\vec{x}_i$ , and the proposed step,  $\vec{y}_{i+1}$ . The proposed step  
1056    is randomly selected from some proposal function  $f(\vec{x}_{i+1}|\vec{x}_i)$ , which depends solely

on the current step (ie. not the further history of the chain). The next step in the chain  $\vec{x}_{i+1}$  can be either the current step or the proposed step determined by whether the proposed step is accepted or rejected. To decide if the proposed step is selected, the acceptance probability,  $\alpha(\vec{x}_i, \vec{y}_i)$ , is calculated as

$$\alpha(\vec{x}_i, \vec{y}_{i+1}) = \min \left( 1, \frac{P(\vec{y}_{i+1}|D)f(\vec{x}_i|\vec{y}_{i+1})}{P(\vec{x}_i|D)f(\vec{y}_{i+1}|\vec{x}_i)} \right). \quad (4.4)$$

Where  $P(\vec{y}_{i+1}|D)$  is the posterior distribution as introduced in section 4.1. To simplify this calculation, the proposal function is required to be symmetric such that  $f(\vec{x}_i|\vec{y}_{i+1}) = f(\vec{y}_{i+1}|\vec{x}_i)$ . In practice, a multi-variate Gaussian distribution is used to throw parameter proposals from. This reduces Equation 4.4 to

$$\alpha(\vec{x}_i, \vec{y}_{i+1}) = \min \left( 1, \frac{P(\vec{y}_{i+1}|D)}{P(\vec{x}_i|D)} \right). \quad (4.5)$$

After calculating this quantity, a random number,  $\beta$ , is generated uniformly between 0 and 1. If  $\beta \leq \alpha(\vec{x}_i, \vec{y}_{i+1})$ , the proposed step is accepted. Otherwise, the chain sets the next step equal to the current step and this procedure is repeated. This can be interpreted as if the posterior probability of the proposed step is greater than that of the current step, ( $P(\vec{y}_{i+1}|D) \geq P(\vec{x}_i|D)$ ), the proposed step will always be accepted. If the opposite is true, ( $P(\vec{y}_{i+1}|D) \leq P(\vec{x}_i|D)$ ), the proposed step will be accepted with probability  $P(\vec{x}_i|D)/P(\vec{y}_{i+1}|D)$ . This ensures that the Markov chain does not get trapped in any local minima in the potentially non-Gaussian posterior distribution. The outcome of this technique is that the density of steps taken in a discrete region is directly proportional to the probability density in that region.

### **1075 4.2.3. MCMC Optimisation**

1076 As discussed in subsection 4.2.2, the proposal function invoked within the MH algo-  
 1077 rithm can take any form and the chain will still converge to the stationary distribution.  
 1078 At each set of proposed parameter values, a prediction of the same spectra has to be  
 1079 generated which requires significant computational resources. Therefore, the number  
 1080 of steps taken before the unique stationary distribution is found should be minimised  
 1081 as only steps after convergence add information to the oscillation analysis. Further-  
 1082 more, the chain should entirely cover the allowable parameter space to ensure that all  
 1083 values have been considered. Tuning the distance that the proposal function jumps  
 1084 between steps on a parameter-by-parameter basis can both minimise the length of the  
 1085 burn-in period and ensure that the correlation between step  $\vec{x}_i$  and  $\vec{x}_j$  is sufficiently  
 1086 small.

1087 The effect of changing the width of the proposal function is highlighted in Figure 4.3.  
 1088 Three scenarios, each with the same underlying stationary distribution (A Gaussian of  
 1089 width 1.0 and mean 0.), are presented. The only difference between the three scenarios  
 1090 is the width of the proposal function, colloquially known as the ‘step size  $\sigma$ ’. Each  
 1091 scenario starts at an initial parameter value of 10.0 which would be considered an  
 1092 extreme variation. For the case where  $\sigma = 0.1$ , it is clear to see that the chain takes  
 1093 a long time to reach the expected region of the parameter. This indicates that this  
 1094 chain would have a large burn-in period and does not converge to the stationary  
 1095 distribution until step  $\sim 500$ . Furthermore, whilst the chain does move towards the  
 1096 expected region, each step is significantly correlated with the previous. Considering  
 1097 the case where  $\sigma = 5.0$ , the chain approaches the expected parameter region almost  
 1098 instantly meaning that the burn-in period is not significant. However, there are clearly  
 1099 large regions of steps where the chain does not move. This is likely due to the chain  
 1100 proposing steps in the tails of the distribution which have a low probability of being

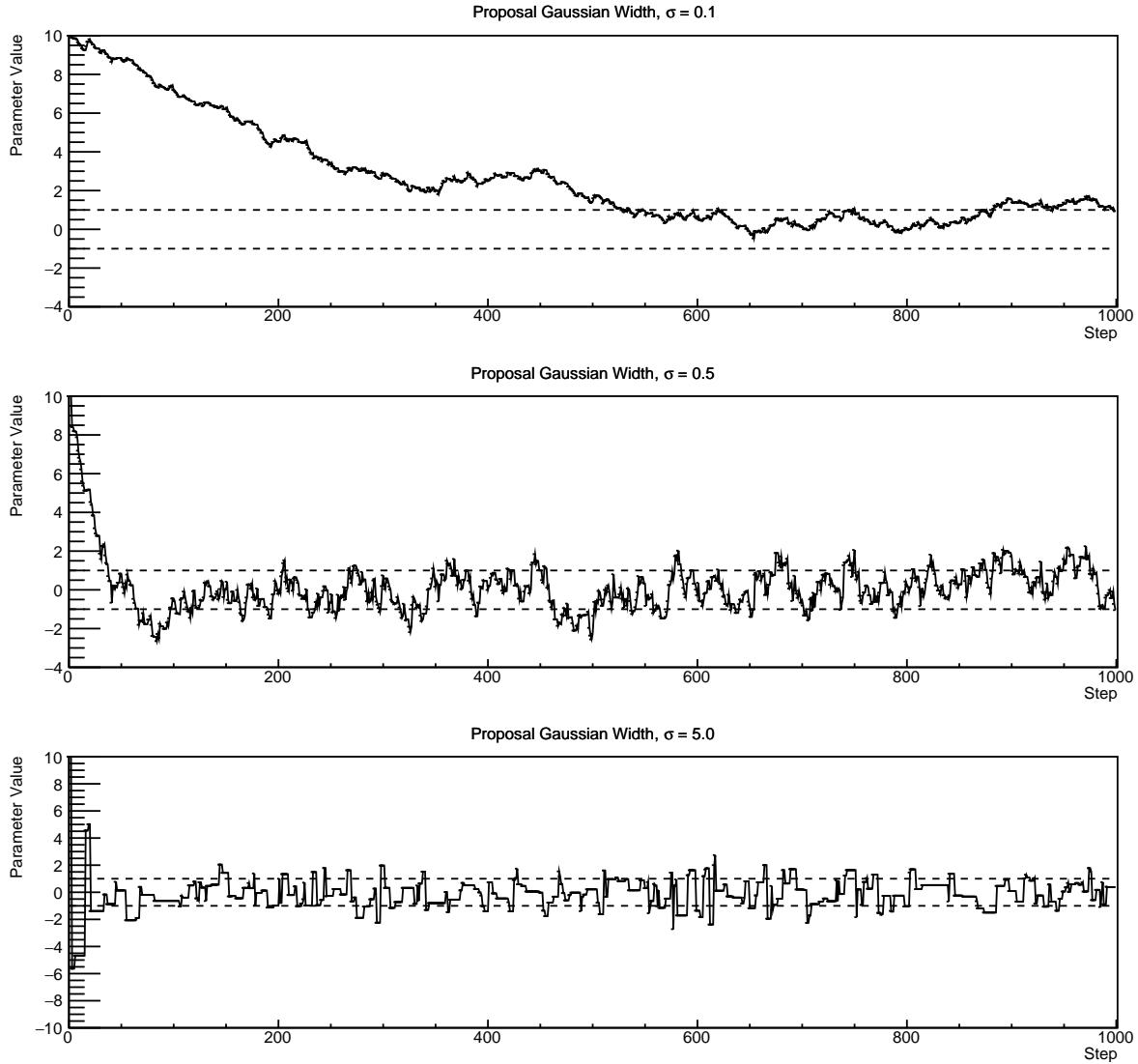
1101 accepted. Consequently, this chain would take a significant number of steps to fully  
 1102 span the allowable parameter region. For the final scenario, where  $\sigma = 0.5$ , you can see  
 1103 a relatively small burn-in period of approximately 100 steps. Once the chain reaches  
 1104 the stationary distribution, it moves throughout the expected region of parameter  
 1105 values many times, sufficiently sampling the full parameter region. This example is a  
 1106 single parameter varying across a continuous distribution and does not fully reflect  
 1107 the difficulties in the many-hundred multi-variate parameter distribution used within  
 1108 this analysis. However, it does give a conceptual idea of the importance of selecting  
 1109 the proposal function and associated step size.

1110 As discussed, step size tuning directly correlates to the average step acceptance  
 1111 rate. If the step size is too small, many steps will be accepted but the chain moves  
 1112 slowly. If the opposite is true, many steps will be rejected as the chain proposes steps  
 1113 in the tails of the distribution. Discussion in [138] suggests that the ‘ideal’ acceptance  
 1114 rate of a high dimension MCMC chain should be approximately  $\sim 25\%$ . An “ideal”  
 1115 step size [138] of

$$\sigma = \frac{2.4}{N_p}, \quad (4.6)$$

1116 where  $N_p$  is the number of parameters included in the MCMC fit. However, the  
 1117 complex correlations between systematics mean that some parameters have to be hand  
 1118 tuned and many efforts have been taken to select a set of parameter-by-parameter step  
 1119 sizes to approximately reach the ideal acceptance rate.

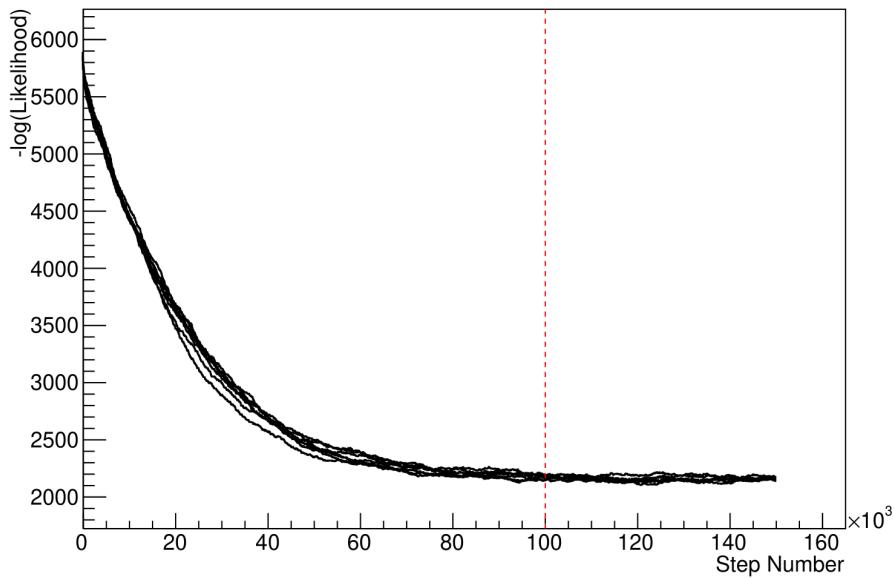
1120 Figure 4.3 highlights the likelihood as calculated by the fit in [DB: Link to AsimovA](#)  
 1121 **Sensitivity Section** as a function of the number of steps in each chain. In practice,  
 1122 many independent MCMC chains are run simultaneously to parallelise the task of



**Figure 4.3.:** Three MCMC chains, each with a stationary distribution equal to a Gaussian centered at 0 and width 1 (As indicated by the black dotted lines). All of the chains use a Gaussian proposal function but have different widths (or ‘step size  $\sigma$ ’). The top panel has  $\sigma = 0.1$ , middle panel has  $\sigma = 0.5$  and the bottom panel has  $\sigma = 5.0$ .

1123 performing the fit. This figure overlays the distribution found in each chain. As seen,  
 1124 the likelihood decreases from its initial value and converges towards a stationary  
 1125 distribution after  $\sim 1 \times 10^5$  steps.

1126 Multiple configurations of this analysis have been performed throughout this thesis  
 1127 where different samples or systematics have been used. For all of these configurations,  
 1128 it was found that a burnin period of  $1 \times 10^5$  was sufficient in all cases.



**Figure 4.4.:** The log-likelihood from the fit detailed in DB: Link to AsimovA Sensitivity Section as a function of the number of steps accumulated in each fit. Many independent MCMC chains were run in parallel and overlaid on this plot. The red line indicates the  $1 \times 10^5$  step burn-in period after which the log-likelihood becomes stable.

### <sup>1129</sup> 4.3. Understanding the MCMC Results

<sup>1130</sup> The previous sections have described how to generate the posterior probability distri-  
<sup>1131</sup> bution using Bayesian MCMC techniques. However, this analysis focuses on oscillation  
<sup>1132</sup> parameter determination. The posterior distribution output from the chain is a high  
<sup>1133</sup> dimension object, with as many dimensions as there are parameters included in the os-  
<sup>1134</sup> cillation analysis. However, this multi-dimensional object is difficult to conceptualize  
<sup>1135</sup> so parameter estimations are often presented in one or two-dimensional projections  
<sup>1136</sup> of this probability distribution. To do this, we invoke the marginalisation technique  
<sup>1137</sup> highlighted in subsection 4.3.1.

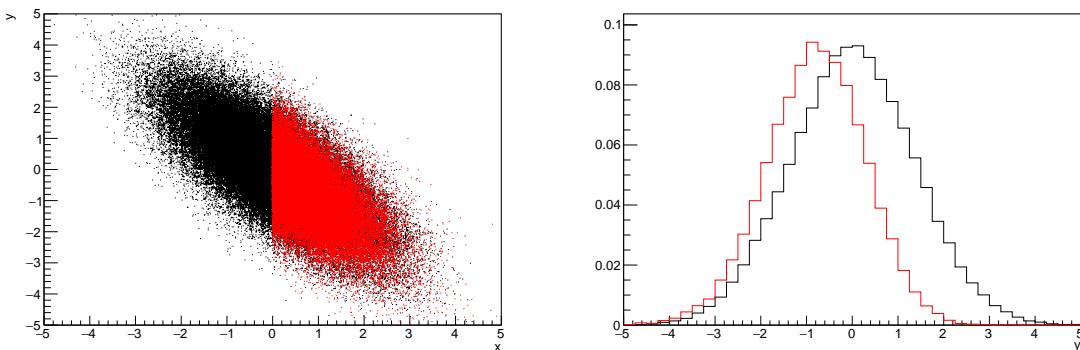
### <sup>1138</sup> 4.3.1. Marginalisation

<sup>1139</sup> The output of the MCMC chain is a highly dimensional probability distribution  
<sup>1140</sup> which is very difficult to interpret. From the standpoint of an oscillation analysis  
<sup>1141</sup> experiment, the one or two-dimensional ‘projections’ of the oscillation parameters of  
<sup>1142</sup> interest are most relevant. Despite this, the best fit values and uncertainties on the  
<sup>1143</sup> oscillation parameters of interest should correctly encapsulate the correlations to the  
<sup>1144</sup> other systematic uncertainties (colloquially called ‘nuisance’ parameters). For this joint  
<sup>1145</sup> beam and atmospheric analysis, the oscillation parameters of interest are  $\sin^2(\theta_{23})$ ,  
<sup>1146</sup>  $\sin^2(\theta_{13})$ ,  $\Delta m_{23}^2$ , and  $\delta_{CP}$ . All other parameters (Including the oscillation parameter  
<sup>1147</sup> this fit is insensitive to) are deemed nuisance parameters. To generate these projections,  
<sup>1148</sup> we rely upon integrating the posterior distribution over all nuisance parameters. This  
<sup>1149</sup> is called marginalisation. A simple example of this technique is to imagine the scenario  
<sup>1150</sup> where two coins are flipped. To determine the probability that the first coin returned  
<sup>1151</sup> a ‘head’, the exact result of the second coin flip is disregarded and simply integrated  
<sup>1152</sup> over. For the parameters of interest,  $\vec{\theta}_i$ , we can calculate the marginalised posterior by  
<sup>1153</sup> integrating over the nuisance parameters,  $\vec{\theta}_n$ . In this case, Equation 4.2 becomes

$$P(\vec{\theta}_i|D) = \frac{\int P(D|\vec{\theta}_i, \vec{\theta}_n)P(\vec{\theta}_i, \vec{\theta}_n)d\vec{\theta}_n}{\int P(D|\vec{\theta})P(\vec{\theta})d\vec{\theta}} \quad (4.7)$$

<sup>1154</sup> Where  $P(\vec{\theta}_i, \vec{\theta}_n)$  encodes the prior knowledge about the uncertainty and correlations  
<sup>1155</sup> between the parameters of interest and the nuisance parameters. In practice, this  
<sup>1156</sup> is simply taking the one or two-dimensional projection of the multi-dimensional  
<sup>1157</sup> probability distribution.

Whilst in principle an easy solution to a complex problem, correlations between the interesting and nuisance parameters can bias the marginalised results. A similar effect is found when the parameters being marginalised over have non-Gaussian probability distributions. For example, Figure 4.5 highlights the marginalisation bias in the probability distribution found for a parameter when requiring a correlated parameter to have a positive parameter value. Due to the complex nature of this oscillation parameter fit presented in this thesis, there are correlations occurring between the oscillation parameters of interest and the other nuisance parameters included in the fit.



**Figure 4.5.:** Left: The two dimensional probability distribution for two correlated parameters  $x$  and  $y$ . The red distribution shows the two dimensional probability distribution when  $0 \leq x \leq 5$ . Right: The marginalised probability distribution for the  $y$  parameter found when requiring the  $x$  to be bound between  $-5 \leq x \leq 5$  and  $0 \leq x \leq 5$  for the black and red distribution, respectively.

### 4.3.2. Parameter Estimation and Credible Intervals

The purpose of this analysis is to determine the best fit values for the oscillation parameters that the beam and atmospheric samples are sensitive to:  $\sin^2(\theta_{23})$ ,  $\sin^2(\theta_{13})$ ,  $\Delta m_{23}^2$ , and  $\delta_{CP}$ . Typically, the results presented take the form of one or two-dimension marginalised probability distributions for the appearance ( $\sin^2(\theta_{13})$  and  $\delta_{CP}$ ) and disappearance ( $\sin^2(\theta_{23})$  and  $\Delta m_{23}^2$ ) parameters. The posterior probability density

<sub>1173</sub> taken from the output MCMC chain is binned in these parameters. The parameter  
<sub>1174</sub> best-fit point is then taken to be the value that has the highest posterior probability.  
<sub>1175</sub> This is performed in both one and two-dimensional projections.

<sub>1176</sub> However, the single best-fit point in a given parameter is not of much use on its  
<sub>1177</sub> own. We would also like to determine the uncertainty, or credible interval, on that  
<sub>1178</sub> best-fit point. The definition of the  $1\sigma$  credible interval is that we have 68% belief that  
<sub>1179</sub> the parameter is within those bounds. For a more generalised definition, the credible  
<sub>1180</sub> interval is the region,  $R$ , of the posterior distribution that contains a specific fraction of  
<sub>1181</sub> the total probability, such that

$$\int_R P(\theta|D)d\theta = \alpha \quad (4.8)$$

<sub>1182</sub> Where  $\theta$  is the parameter on which we calculate the credible interval. This technique  
<sub>1183</sub> then calculates the  $\alpha \times 100\%$  credible interval.

<sub>1184</sub> In practice, this analysis uses the highest posterior density (HPD) credible intervals  
<sub>1185</sub> which are calculated through the following method. First, the probability distribution  
<sub>1186</sub> is area-normalised such that it has an integrated area equal to 1.0. The bins of proba-  
<sub>1187</sub> bility are then summed from the highest to lowest until the sum exceeds the  $1\sigma$  level  
<sub>1188</sub> (0.68 in this example). This process is repeated for a range of credible intervals, notably  
<sub>1189</sub> the  $1\sigma$ ,  $2\sigma$  and  $3\sigma$  along with other levels where the critical values for each level can  
<sub>1190</sub> be found in [139]. This process can be repeated for the two-dimensional probability  
<sub>1191</sub> distributions by creating two-dimensional contours of credible intervals rather than a  
<sub>1192</sub> one-dimensional result.

---

### <sup>1193</sup> 4.3.3. Bayesian Model Comparisons

<sup>1194</sup> Due to the matter resonance, this analysis has some sensitivity to the mass hierarchy  
<sup>1195</sup> of neutrino states (whether  $\Delta m_{23}^2$  is positive or negative) and the octant of  $\sin^2(\theta_{23})$   
<sup>1196</sup> . The Bayesian approach utilised within this analysis gives an intuitive method of  
<sup>1197</sup> model comparison by determining which hypothesis is most favourable. Taking the  
<sup>1198</sup> ratio of Equation 4.3 for the two hypotheses of normal hierarchy,  $NH$ , and inverted  
<sup>1199</sup> hierarchy,  $IH$ , gives

$$\frac{P(\vec{\theta}_{NH}|D)}{P(\vec{\theta}_{IH}|D)} = \frac{P(D|\vec{\theta}_{NH})}{P(D|\vec{\theta}_{IH})} \times \frac{P(\vec{\theta}_{NH})}{P(\vec{\theta}_{IH})}. \quad (4.9)$$

<sup>1200</sup> The middle term defines the Bayes factor which is a data-driven interpretation of  
<sup>1201</sup> how strong the data prefers one hierarchy to the other. For this analysis, equal priors  
<sup>1202</sup> on both mass hierarchy hypotheses are chosen ( $P(\vec{\theta}_{NH}) = P(\vec{\theta}_{IH}) = 0.5$ ). In practice,  
<sup>1203</sup> the MCMC chain proposes a value of  $|\Delta m_{23}^2|$  and then applies a 50% probability  
<sup>1204</sup> that the value is sign flipped. Consequently, the Bayes factor can be calculated from  
<sup>1205</sup> the ratio of the probability density in either hypothesis. This equates to counting the  
<sup>1206</sup> number of steps taken in the normal and inverted hierarchies and taking the ratio. The  
<sup>1207</sup> same approach can be taken to compare the upper octant (UO) compared to the lower  
<sup>1208</sup> octant (LO) hypothesis of  $\sin^2(\theta_{23})$ .

<sup>1209</sup> Whilst the value of the Bayes factor should always be shown, the Jeffreys scale [140]  
<sup>1210</sup> (highlighted in Table 4.1) gives an indication of the strength of preference for one model  
<sup>1211</sup> compared to the other. Other interpretations of the strength of preference of a model  
<sup>1212</sup> exist, e.g. the Kass and Raferty Scale [141].

$\log_{10}(B_{AB})$	$B_{AB}$	Strength of Preference
< 0.0	< 1	No preference for hypothesis A (Supports hypothesis B)
0.0 – 0.5	1.0 – 3.16	Preference for hypothesis A is weak
0.5 – 1.0	3.16 – 10.0	Preference for hypothesis A is substantial
1.0 – 1.5	10.0 – 31.6	Preference for hypothesis A is strong
1.5 – 2.0	31.6 – 100.0	Preference for hypothesis A is very strong
> 2.0	> 100.0	Decisive preference for hypothesis A

**Table 4.1.:** Jeffreys scale for strength of preference for two models  $A$  and  $B$  as a function of the calculated Bayes factor ( $B_{AB} = B(A/B)$ ) between the two models [140]. The original scale is given in terms of  $\log_{10}(B(A/B))$  but converted to linear scale for easy comparison throughout this thesis.

#### <sup>1213</sup> 4.3.4. Comparison of MCMC Output to Expectation

<sup>1214</sup> To ensure the fit is performing well, a best-fit spectrum is produced using the pos-  
<sup>1215</sup> terior probability distribution and compared with the data, allowing easy by-eye  
<sup>1216</sup> comparisons to be made. A simple method of doing this is to perform a comparison  
<sup>1217</sup> in the fitting parameters (For instance, the reconstructed neutrino energy and lepton  
<sup>1218</sup> direction for T2K far detector beam samples) of the spectra generated by the MCMC  
<sup>1219</sup> chain to ‘data’. This ‘data’ could be true data or some variation of Monte Carlo predic-  
<sup>1220</sup> tion. This allows easy comparison of the MCMC probability distribution to the data.  
<sup>1221</sup> To perform this,  $N$  steps from the post burn-in MCMC chain are randomly selected  
<sup>1222</sup> (Where for all plots of this style in this thesis,  $N = 3000$ ). From these, the Monte Carlo  
<sup>1223</sup> prediction at each step is generated by reweighting the model parameters to the values  
<sup>1224</sup> specified at that step. Due to the probability density being directly correlated with  
<sup>1225</sup> the density of steps in a certain region, parameter values close to the best fit value are  
<sup>1226</sup> most likely to be selected.

<sup>1227</sup> In practice, for each bin of the fitting parameters has a probability distribution  
<sup>1228</sup> of event rates, with one entry per sampled MCMC step. This distribution is binned  
<sup>1229</sup> where the bin with the highest probability is selected as the mean and an error on

1230 the width of this probability distribution is calculated using the approach highlighted  
1231 in subsection 4.3.2. Consequently, the best fit distribution in the fit parameter is not  
1232 necessarily that which would be attained by reweighting the Monte Carlo prediction  
1233 to the most probable parameter values.

1234 A similar study can be performed to illustrate the freedom of the model parameter  
1235 space prior to the fit. This can be done by throwing parameter values from the prior  
1236 uncertainty of each parameter. This becomes troublesome for parameters with no  
1237 prior uncertainty as the range is technically infinite. Where applicable solutions to  
1238 remove these have been addressed.

1239 **Chapter 5**

1240 **Simulation, Reconstruction, and Event  
Reduction**

1242 As a crucial part of the oscillation analysis, an accurate prediction of the expected  
1243 neutrino spectrum at the far detector is required. This includes modeling the flux  
1244 generation, neutrino interactions, and detector effects. All of the simulation packages  
1245 required to do this are briefly described in section 5.1. The reconstruction of neutrino  
1246 events inside the far detector, including the `fitQun` algorithm, is documented in  
1247 section 5.2. This also includes data quality checks of the SK-V data which the author  
1248 performed for the T2K oscillation analysis presented at Neutrino 2020 [80]. Finally,  
1249 section 5.3 describes the steps taken in the SK detector to trigger on events of interest  
1250 whilst removing the comparatively large rate of cosmic ray muon events.

1251 **5.1. Simulation**

1252 In order to generate a Monte Carlo prediction of the expected event rate at the far  
1253 detector, all the processes in the beam and atmospheric flux, neutrino interaction, and  
1254 detector need to be modeled. Each of these parts is individually modeled and each of  
1255 them is detailed below.

1256 The beamline simulation consists of three distinct parts: the initial hadron inter-  
1257 action modeled by FLUKA [142], the target station geometry and particle tracking

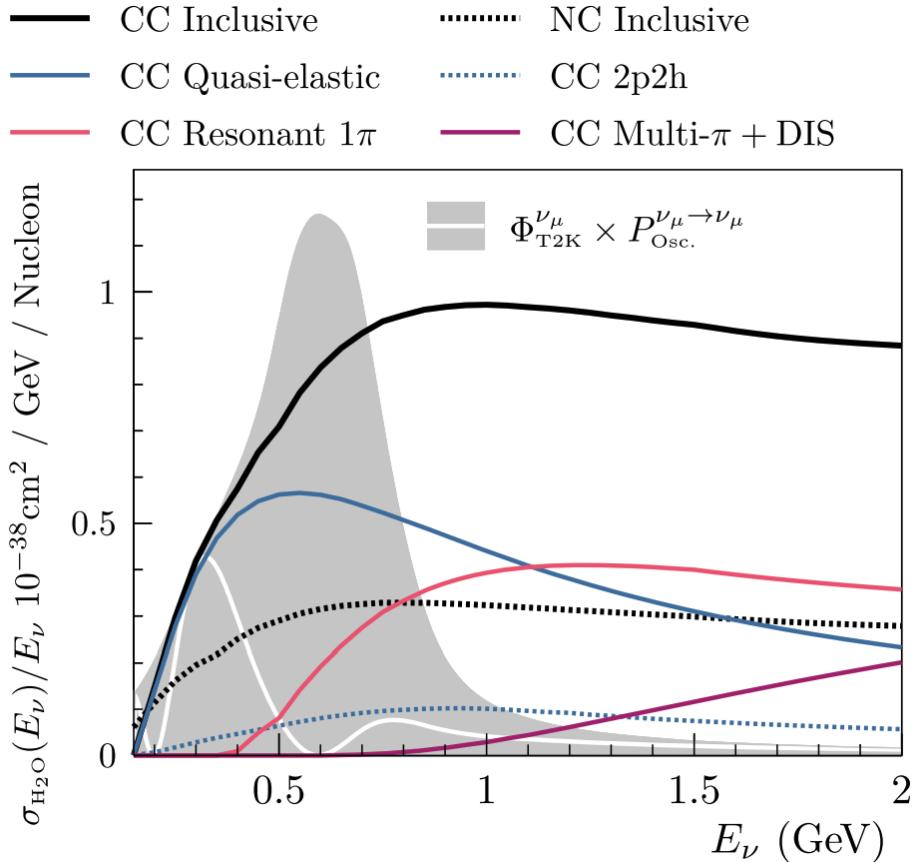
performed by JNUBEAM, [143, 144] and any hadronic re-interactions simulated by GCALOR [145]. The primary hadronic interactions are  $O(10)\text{GeV}$ , where FLUKA matches external cross-section data better than GCALOR [146]. However, FLUKA is not very adaptable so a small simulation is built to model the interactions in the target and the output is then passed to JNUBEAM and GCALOR for propagation. The hadronic interactions are tuned to data from the NA61/SHINE [147–149] and HARP [150] experiments. The tuning is done by reweighting the FLUKA and GCALOR predictions to match the external data multiplicity and cross-section measurements, based on final state particle kinematics [146]. The culmination of this simulation package generates the predicted flux for neutrino and antineutrino beam modes which are illustrated in Figure 3.7.

The atmospheric neutrino flux is simulated by the HKKM model [43, 45]. The primary cosmic ray flux is tuned to AMS [151] and BESS [152] data assuming the US-standard atmosphere '76 [153] density profile and includes geomagnetic field effects. The primary cosmic rays interact to generate pions and muons. The interaction of these secondary particles to generate neutrinos is handled by DPMJET-III [154] for energies above 32GeV and JAM [45, 155] for energies below that value **DB: Question for Giles: Why different generators for above/below 32GeV?** These hadronic interactions are tuned to BESS and L3 data [156, 157] using the same methodology as the tuning of the beamline simulation. The energy and cosine zenith predictions of  $\nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$  flux are given in Figure 2.3 and Figure 2.5, respectively. The flux is approximately symmetrical and peaked around the horizon ( $\cos(\theta_Z) = 0.0$ ). This is because horizontally-going pions and kaons can travel further than their vertically-going counterparts resulting in a larger probability of decaying to neutrinos. The symmetry is broken in lower-energy neutrinos due to geomagnetic effects, which modify the track of the primary cosmic rays. Updates to the HKKM model are currently ongoing [158].

Once a flux prediction has been made for all three detectors, NEUT 5.4.0 [159, 160] models the interactions of the neutrinos in the detectors. For the purposes of this analysis, quasi-elastic (QE), meson exchange (MEC), single meson production (PROD), coherent pion production (COH), and deep inelastic scattering (DIS) interactions are simulated. These interaction categories can be further broken down by whether they were propagated via a  $W^\pm$  boson in Charged Current (CC) interactions or via a  $Z^0$  boson in Neutral Current (NC) interactions. CC interactions have a charged lepton in the final state, which can be flavour-tagged in reconstruction to determine the flavour of the neutrino. In contrast, NC interactions have a neutrino in the final state so no flavour information can be determined from the observables left in the detector after an interaction. This is the reason why neutrinos which interact through NC modes are assumed to not oscillate within this analysis. Both CC and NC interactions are modeled for all the above interaction categories, other than MEC interactions which are only modeled for CC events.

As illustrated in Figure 5.1, CCQE interactions dominate the cross-section of neutrino interactions around  $E_\nu \sim 0.5\text{GeV}$ . The NEUT implementation adopts the Llewellyn Smith [161] model for neutrino-nucleus interactions, where the nuclear ground state of any bound nucleons (neutrino-oxygen interactions) is approximated by a spectral-function [162] model that simulates the effects of Fermi momentum and Pauli blocking. The cross-section of QE interactions is controlled by vector and axial-vector form factors parameterised by the BBBA05 [163] model and a dipole form factor with  $M_A^{QE} = 1.21\text{GeV}$  fit to external data [164], respectively. NEUT implements the Valencia [165] model to simulate MEC events, where two nucleons and two holes in the nuclear target are produced (Often called 2p2h interactions).

For neutrinos of energy  $O(1)\text{GeV}$ , PROD interactions become dominant. These predominantly produce charged and neutral pions although  $\gamma$ , kaon, and  $\eta$  production

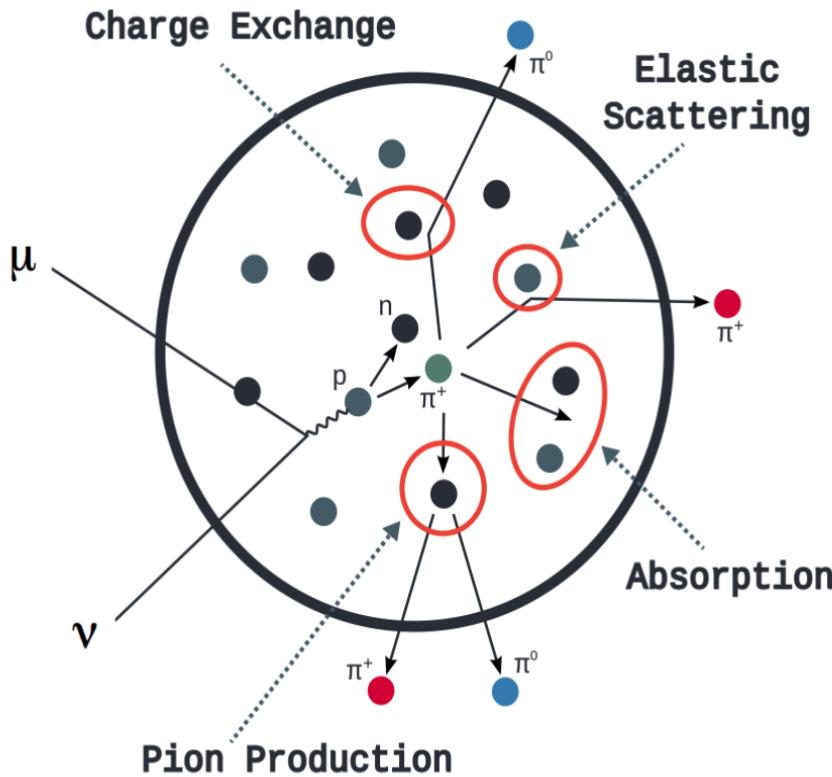


**Figure 5.1.:** The NEUT prediction of the  $\nu_\mu$ -H<sub>2</sub>O cross-section overlaid on the T2K  $\nu_\mu$  flux. The charged current (black, solid) and neutral current (black, dashed) inclusive, charged current quasi-elastic (blue, solid), charged current 2p2h (blue, dashed), charged current single pion production (pink), and charged current multi- $\pi$  and DIS (Purple) cross-sections are illustrated. Figure taken from [159].

is also considered. To simulate these interactions, the Berger-Sehgal [166] model is implemented within NEUT. It simulates the excitation of a nucleon from a neutrino interaction, production of an intermediate baryon, and the subsequent decay to a single meson or  $\gamma$ . Pions can also be produced through COH interactions, which occur when the incoming neutrino interacts with the entire oxygen nucleus leaving a single pion outside of the nucleus. NEUT utilises the Berger-Sehgal [167] model to simulate these COH interactions.

DIS and multi- $\pi$  producing interactions become the most dominant for energies  $> O(5)\text{GeV}$ . PYTHIA [168] is used to simulate any interaction with invariant mass

<sub>1319</sub>  $W > 2\text{GeV}/c^2$ , which produces at least one meson. For any interaction which produces  
<sub>1320</sub> at least two mesons but has  $W < 2\text{GeV}/c^2$ , the Bronner model is used [169]. Both  
<sub>1321</sub> of these models use parton distribution functions based on the Bodek-Yang model  
<sub>1322</sub> [170–172].



**Figure 5.2.:** Illustration of the various processes which a pion can undergo before exiting the nucleus. Taken from [173].

<sub>1323</sub> Any pion which is produced within the nucleus can re-interact through final state  
<sub>1324</sub> interactions before it exits, as illustrated by the scattering, absorption, production, and  
<sub>1325</sub> exchange interactions in Figure 5.2. These re-interactions alter the observable particles  
<sub>1326</sub> within the detector. For instance, if the charged pion from a CC PROD interaction is  
<sub>1327</sub> absorbed, the observables would mimic a CC QE interaction. To simulate these effects,  
<sub>1328</sub> NEUT uses a semi-classical intranuclear cascade model [159]. This cascade functions by  
<sub>1329</sub> stepping the pion through the nucleus in fixed-length steps equivalent to  $dx = R_N/100$ ,

1330 where  $R_N$  is the radius of the nucleus. At each step, the simulation allows the pion  
1331 to interact through scattering, charged exchange, absorption, or production with an  
1332 interaction-dependent probability calculated from a fit to external data [174]. This  
1333 cascade continues until the pion is absorbed or exits the nucleus.

1334 Once the final state particle kinematics have been determined by NEUT, they are  
1335 passed into the detector simulation. The near detectors, ND280 and INGRID, are  
1336 simulated using a GEANT4 package [118,175] to simulate the detector geometry, particle  
1337 tracking, and energy deposition. The response of the detectors is simulated using  
1338 the elecSim package [118]. The far detector simulation is based upon the original  
1339 Kamiokande experiment software which uses the GEANT3-based SKDETSIM [118,176]  
1340 package. This simulates the interactions of particles in the water as well as Cherenkov  
1341 light production. The water quality and PMT calibration measurements detailed in  
1342 subsection 3.1.2 are also used within this simulation to make accurate predictions of  
1343 the detector response.

## 1344 5.2. Event Reconstruction at SK

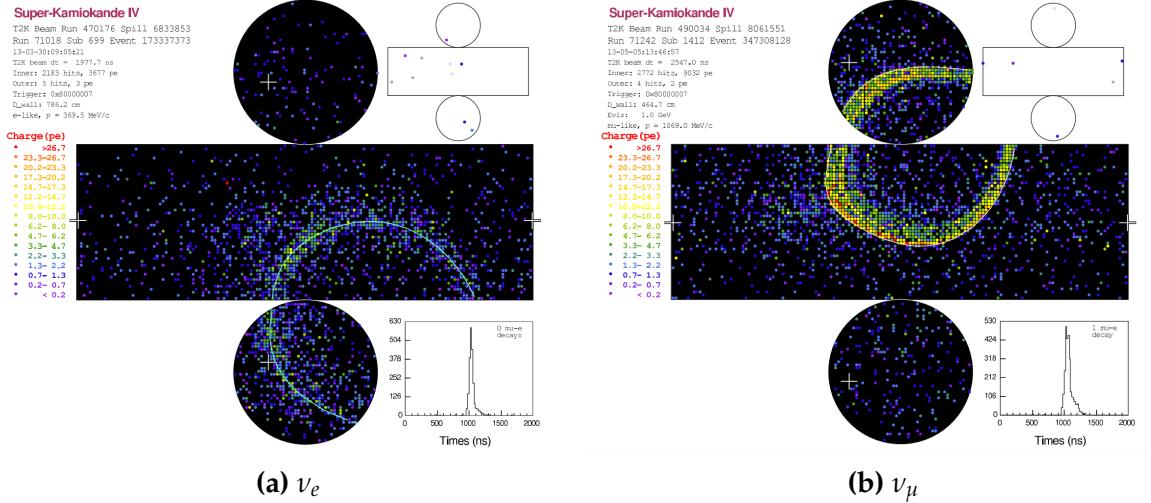
1345 Any event which generates optical photons that occurs in SK will be observed by the  
1346 PMT array, where each PMT records the time and accumulated charge. This recorded  
1347 information is shown in event displays similar to those illustrated in Figure 5.3. To  
1348 be useful for physics analyses, this series of PMT hit information needs to be recon-  
1349 structed to determine the particle's identity and kinematics (or track parameters):  
1350 four-vertex, direction, and momenta. The reconstruction uses the fact that the charge  
1351 and timing distribution of photons generated by a particular particle in an event is  
1352 dependent upon its initial kinematics. Electron and muon rings are distinguished by  
1353 their "fuzziness". Muons are heavier and less affected by scattering or showering

1354 meaning they typically produce “crisp” rings. Electrons are more likely to interact  
1355 via electromagnetic showering or scattering which results in larger variations of their  
1356 direction from the initial direction. Consequently, electrons typically produce “fuzzier”  
1357 rings compared to muons.

1358 For the purposes of this analysis, the `fitQun` reconstruction algorithm is utilised.  
1359 Its core function is to compare a prediction of the accumulated charged and timing  
1360 distribution from each PMT, generated for a particular particle identity and track  
1361 parameters, to that observed in the neutrino event. It determines the preferred values  
1362 by minimising a likelihood function which includes information from PMTs which  
1363 were hit and those that were not hit. `fitQun` performs a simultaneous fit of particle  
1364 kinematics and identity, improving both the accuracy of the fit parameters and the  
1365 rejection of neutral current  $\pi^0$  events [177, 178]. The `fitQun` algorithm is based on  
1366 the key concepts of the MiniBooNE reconstruction algorithm [179] and is described  
1367 in [180] which is summarised below. The `fitQun` algorithm improves upon the APFit  
1368 reconstruction algorithm which has been used for many previous SK analyses. APFit  
1369 fits the vertex from timing information and then fits the momentum and direction  
1370 of the particle from PMT hits within a 43 deg Cherenkov cone (which assumes an  
1371 ultra-relativistic particle). It then fits the particle identity once the track parameters  
1372 have been fit.

1373 The `fitQun` reconstruction algorithm proceeds by:

- 1374 • **Vertex pre-fitting:** An estimate of the vertex is made using a goodness-of-fit  
1375 metric
- 1376 • **Peak finding:** The initial time of the event is determined by clustering events by  
1377 time residuals



**Figure 5.3.:** Event displays from Super Kamiokande illustrating the “crisp” ring from a muon and the typically “fuzzy” electron ring. Each pixel represents a PMT and the color scheme denotes the accumulated charge deposited on that PMT. Figures taken from [181].

- **Single-ring fits:** Given the pre-fit vertex and estimated time of interaction, a maximum likelihood technique searches for a single particle generating light
  - **Multi-ring fits:** Seeded from the single-ring fits, hypotheses with multiple light-producing particles are considered using the same maximum likelihood technique
- An event in SK can consist of a primary and decay particles. For example, a charged current muon neutrino interaction can generate two particles that have the potential of generating Cherenkov photons: the primary muon, and the secondary decay-electron from the muon. To ensure the particles are reconstructed separately, each event is divided into time clusters which are called “subevents”. Subevents after the primary subevent are considered to be decay electrons. To find all the subevents in an event, a vertex goodness metric is calculated for some vertex position  $\vec{x}$  and time  $t$ ,

$$G(\vec{x}, t) = \sum_i^{\text{hit PMTs}} \exp \left( -\frac{1}{2} \left( \frac{T_{\text{Res}}^i(\vec{x}, t)}{\sigma} \right)^2 \right) \quad (5.1)$$

1389 where

$$T_{Res}^i(\vec{x}, t) = t^i - t - |R_{PMT}^i - \vec{x}| / c_n \quad (5.2)$$

1390 is the residual hit time. It is the difference in time between the PMT hit time,  $t^i$ ,  
1391 of the  $i^{th}$  PMT and the expected time of the PMT hit if the photon was emitted at  
1392 the start of the vertex.  $R_{PMT}^i$  is the position of the  $i^{th}$  PMT,  $c_n$  is the speed of light in  
1393 water and  $\sigma = 4\text{ns}$  which is comparable to the time resolution of the PMT. When the  
1394 proposed fit values of time and vertex are close to the true values,  $T_{Res}^i(\vec{x}, t)$  tends to  
1395 zero resulting in subevents appearing as spikes in the goodness metric. The proposed  
1396 fit vertex and time are grid-scanned, and the values which maximise the goodness  
1397 metric are selected as the “pre-fit vertex”. Whilst this predicts a vertex for use in  
1398 the clustering algorithm, the final vertex is fit using the higher-precision maximum  
1399 likelihood method described below.

1400 Once the pre-fit vertex has been determined, the goodness metric is scanned as  
1401 a function of  $t$  to determine the number of subevents. A peak-finding algorithm is  
1402 then used on the goodness metric, requiring the goodness metric to exceed some  
1403 threshold and drop below a reduced threshold before any subsequent additional  
1404 peaks are considered. The thresholds are set such that the rate of false peak finding  
1405 is minimised while still attaining good data to Monte Carlo agreement. To improve  
1406 performance, the pre-fit vertex for each delayed subevent is re-calculated after PMT  
1407 hits from the previous subevent are masked. This improves the decay-electron tagging  
1408 performance. Once all subevents have been determined, the time window around  
1409 each subevent is then defined by the earliest and latest time which satisfies  $-180 <$

<sub>1410</sub>  $T_{Res}^i < 800\text{ns}$ . The subevents and associated time windows are then used as seeds for  
<sub>1411</sub> further reconstruction.

<sub>1412</sub> For a given subevent, the `fitQun` algorithm constructs a likelihood based on the  
<sub>1413</sub> accumulated charge  $q_i$  and time information  $t_i$  from the  $i^{th}$  PMT,

$$L(\Gamma, \vec{\theta}) = \prod_j^{\text{unhit}} P_j(\text{unhit}|\Gamma, \vec{\theta}) \prod_i^{\text{hit}} \{1 - P_i(\text{unhit}|\Gamma, \vec{\theta})\} f_q(q_i|\Gamma, \vec{\theta}) f_t(t_i|\Gamma, \vec{\theta}), \quad (5.3)$$

<sub>1414</sub> where  $\vec{\theta}$  defines the track parameters; vertex position, direction vector and mo-  
<sub>1415</sub> ments, and  $\Gamma$  represents the particle hypothesis.  $P_i(\text{unhit}|\Gamma, \vec{\theta})$  is the probability of the  
<sub>1416</sub>  $i^{th}$  tube to not register a hit given the track parameters and particle hypothesis. The  
<sub>1417</sub> charge likelihood,  $f_q(q_i|\Gamma, \vec{\theta})$ , and time likelihood,  $f_t(t_i|\Gamma, \vec{\theta})$ , represents the probability  
<sub>1418</sub> density function of observing charge  $q_i$  and time  $t_i$  on the  $i^{th}$  PMT given the specified  
<sub>1419</sub> track parameters and particle hypothesis.

<sub>1420</sub> The predicted charge is calculated based on contributions from both the direct  
<sub>1421</sub> light and the scattered light. The direct light contribution is determined based on the  
<sub>1422</sub> integration of the Cherenkov photon profile along the track. PMT angular acceptance,  
<sub>1423</sub> water quality, and calibration measurements discussed in subsection 3.1.2 are included  
<sub>1424</sub> to accurately predict the charge probability density at each PMT. The scattered light  
<sub>1425</sub> is calculated in a similar way, although it includes a scattering function that depends  
<sub>1426</sub> on the vertex of the particle and the position of the PMT. The charge likelihood is  
<sub>1427</sub> calculated by comparing the prediction to the observed charge in the PMT.

<sub>1428</sub> The time likelihood is approximated to depend on the vertex  $\vec{x}$ , direction  $\vec{d}$ , and  
<sub>1429</sub> time  $t$  of the track as well as the particle hypothesis. The expected time for PMT hits is

<sup>1430</sup> calculated by assuming unscattered photons being emitted from the midpoint of the  
<sup>1431</sup> track,  $S_{mid}$ ,

$$t_{exp}^i = t + S_{mid}/c + |R_{PMT}^i - \vec{x} - S_{mid}\vec{d}|/c_n, \quad (5.4)$$

<sup>1432</sup> where  $c$  is the speed of light in a vacuum. The time likelihood is then expressed in  
<sup>1433</sup> terms of the residual difference between the PMT hit time and the expected hit time,  
<sup>1434</sup>  $t_{Res}^i = t^i - t_{exp}^i$ . The particle hypothesis and momentum also affect the Cherenkov  
<sup>1435</sup> photon distribution. These parameters modify the shape of the time likelihood density  
<sup>1436</sup> since in reality not all photons are emitted at the midpoint of the track. As with the  
<sup>1437</sup> charge likelihood, the contributions from both the direct and scattered light to the time  
<sup>1438</sup> likelihood density are calculated separately, which are both calculated from particle  
<sup>1439</sup> gun studies.

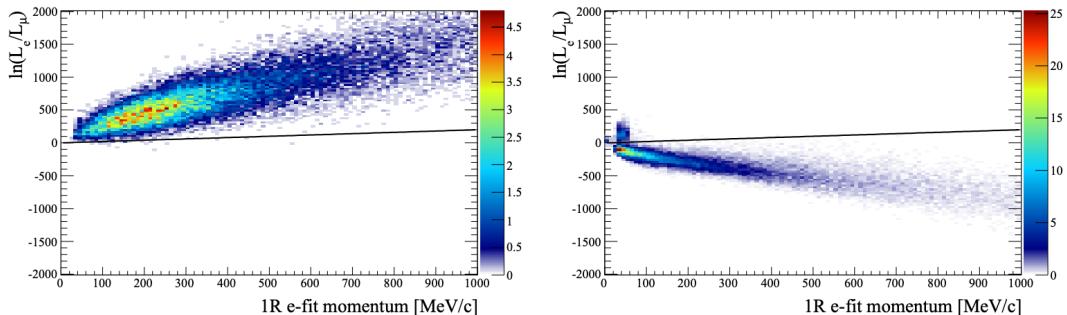
<sup>1440</sup> The track parameters and particle identity which maximise  $L(\Gamma, \vec{\theta})$  are defined as  
<sup>1441</sup> the best-fit parameters. In practice MINUIT [182] is used to minimise the value of  
<sup>1442</sup>  $-\ln L(\Gamma, \vec{\theta})$ . The `fitQun` algorithm considers an electron-like, muon-like, and charged  
<sup>1443</sup> pion-like hypothesis for events with a single final state particle, denoted “single-ring  
<sup>1444</sup> events”. The particle’s identity is determined by taking the ratio of the likelihood  
<sup>1445</sup> of each of the hypotheses. For instance, electrons and muons are distinguished by  
<sup>1446</sup> considering the value of  $\ln(L(e, \vec{\theta}_e)/L(\mu, \vec{\theta}_\mu))$  in comparison to the reconstructed  
<sup>1447</sup> momentum of the electron hypothesis, as illustrated by Figure 5.4. The coefficients of  
<sup>1448</sup> the discriminator between electron-like and muon-like events are determined from  
<sup>1449</sup> Monte Carlo studies [180]. Similar distributions exist for distinguishing electron-like  
<sup>1450</sup> events from  $\pi^0$ -like events, and muon-like events from pion-like events. The cuts are  
<sup>1451</sup> defined as,

$$\text{Electron/Muon} : \ln(L_e/L_\mu) > 0.2 \times p_e^{rec} [\text{MeV}],$$

$$\text{Electron}/\pi^0 : \ln(L_e/L_{\pi^0}) < 175 - 0.875 \times m_{\gamma\gamma} [\text{MeV}], \quad (5.5)$$

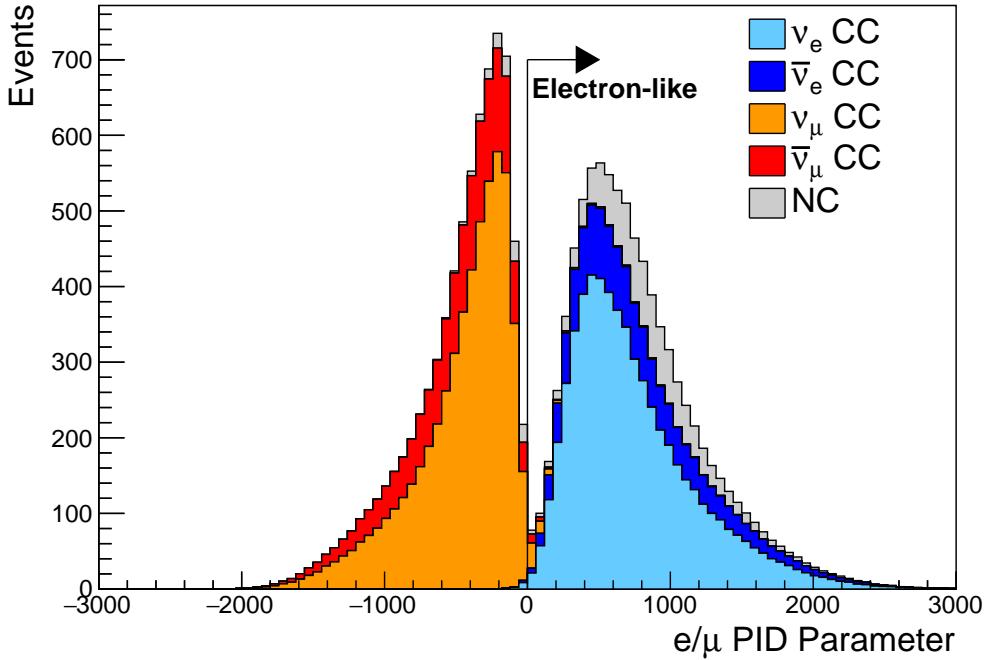
$$\text{Muon/Pion} : \ln(L_\mu/L_{\pi^\pm}) < 0.15 \times p_\mu^{rec} [\text{MeV}],$$

as taken from [183], where  $p_e^{rec}$  and  $p_\mu^{rec}$  are the reconstructed momentum of the single-ring electron and muon fits, respectively.  $m_{\gamma\gamma}$  represents the reconstructed invariant mass of the two photons emitted from  $\pi^0$  decay. Typically, the distance between a particular entry in these two-dimensional distributions and the cut-line is termed the PID parameter and is illustrated in Figure 5.5.



**Figure 5.4.:** The difference of the electron-like and muon-like log-likelihood compared to the reconstructed single-ring fit momentum for atmospheric  $\nu_e$  (left) and  $\nu_\mu$  (right) samples. The black line represents the cut used to discriminate electron-like and muon-like events, which coefficients obtained from Monte Carlo studies. Figures taken from [180].

The `fitQun` algorithm also considers a  $\pi^0$  hypothesis. To do this, it performs a fit looking for two standard electron-hypothesis tracks which point to the same four-vertex. This assumes the electron tracks are generated from photon-conversion so the electron tracks actually appear offset from the proposed  $\pi^0$  vertex. For these fits, the conversion length, direction, and momentum of each photon are also considered as track parameters which are then fit in the same methodology as the standard single-ring hypotheses.



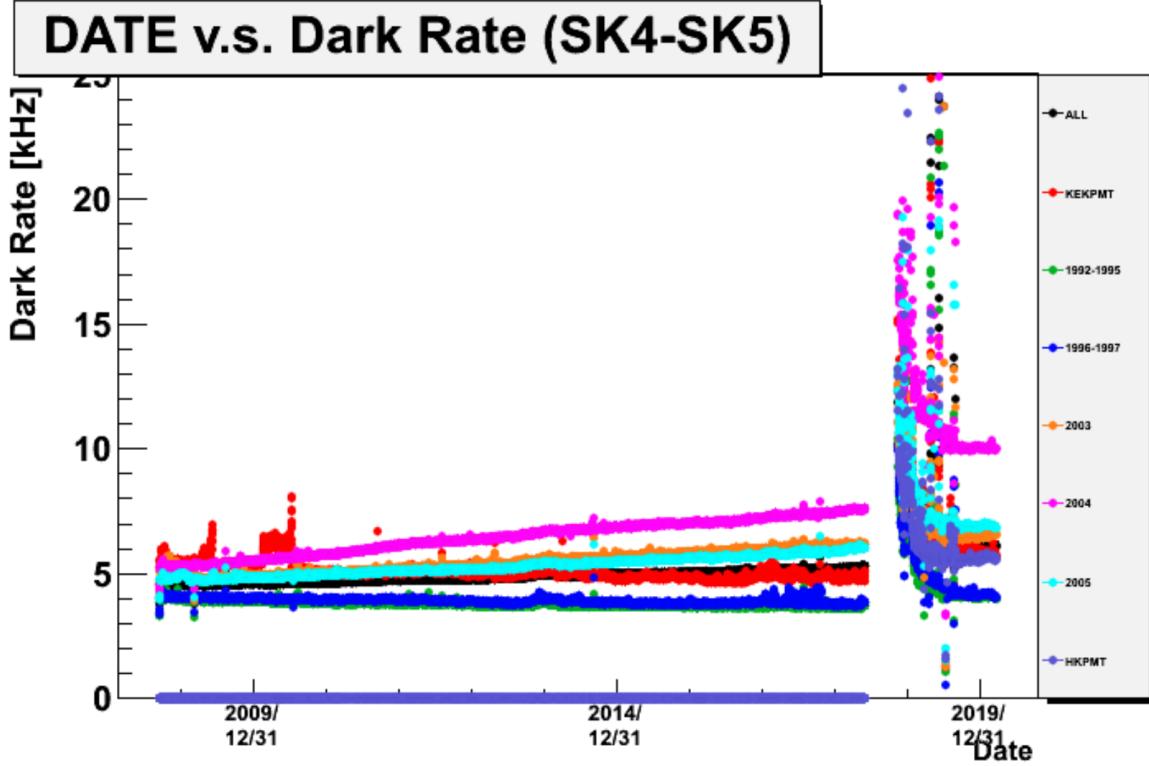
**Figure 5.5.:** The electron/muon PID separation parameter for all sub-GeV single-ring events in SK-IV. The charged current (CC) component is broken down in four flavours of neutrino ( $\nu_\mu$ ,  $\bar{\nu}_\mu$ ,  $\nu_e$  and  $\bar{\nu}_e$ ). Events with positive values of the parameter are determined to be electron-like.

Whilst lower energy events are predominantly single-ring events, higher energy neutrino events can generate final states with multiple particles which generate Cherenkov photons. These “multi-ring” hypotheses are also considered in the `fitQun` algorithm. When calculating the charge likelihood density, the predicted charge associated with each ring is calculated separately and then summed to calculate the total accumulated charge on each PMT. Similarly, the time likelihood for the multi-ring hypothesis is calculated assuming each ring is independent. Each track is time-ordered based on the time of flight from the center of the track to the PMT and the direct light from any ring incident on the PMT is assumed to arrive before any scattered light. To reduce computational resource usage, the multi-ring fits only consider electron-like and charged pion-like rings as the pion fit can be used as a proxy for a muon fit due to their similar mass.

1476 Multi-ring fits proceed by proposing another ring to the previous fit and then  
1477 fitting the parameters in the method described above. Typically, multi-ring fits have  
1478 the largest likelihood because of the additional degrees of freedom introduced. A  
1479 likelihood value is calculated for the  $n$ -ring and  $(n + 1)$ -ring hypotheses, where the  
1480 additional ring is only included if the likelihood value is above 9.35, based on Monte  
1481 Carlo studies in [184].

### 1482 5.2.1. Validation of Reconstruction in SK-V

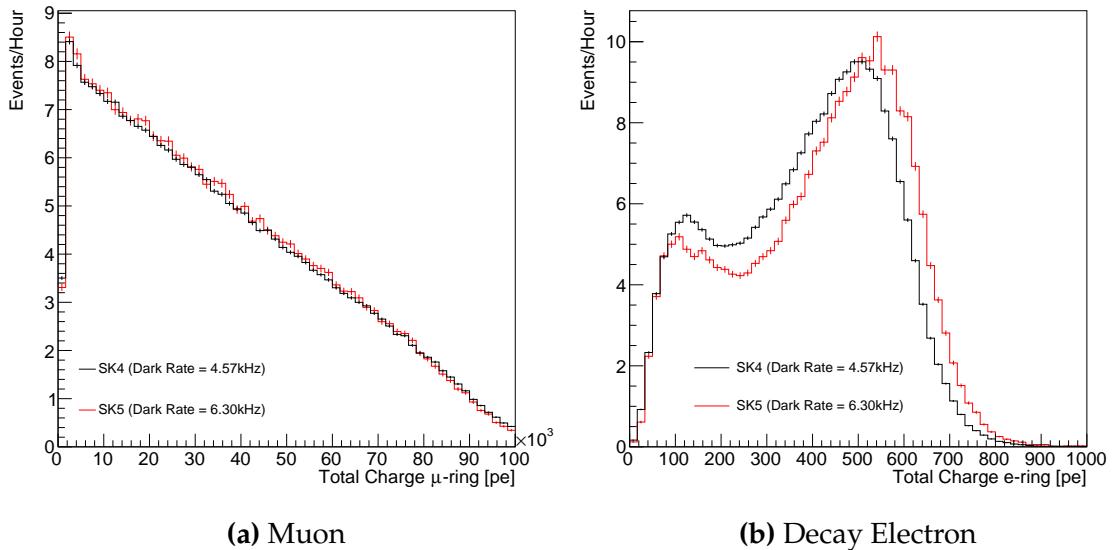
1483 As an example of how the reconstruction depends on the detector conditions, the  
1484 author of this thesis assessed the quality of event reconstruction for SK-V data. The  
1485 detector systematics used in the T2K-only oscillation analysis are determined using  
1486 data-to-Monte Carlo comparisons of the SK-IV data [185]. Due to tank-open mainte-  
1487 nance occurring between SK-IV and SK-V, the dark rate of each PMT was observed to  
1488 increase in SK-V due to light exposure for a significant time during the repairs. This  
1489 increase can be seen in Figure 5.6. Run-10 of the T2K experiment was conducted in  
1490 the SK-V period, so the consistency of SK-IV and SK-V data needs to be studied to  
1491 determine whether the SK-IV-defined systematics can be applied to the run-10 data.  
1492 This comparison study was performed using the stopping muon data set for both the  
1493 SK-IV and SK-V periods. This data sample is used due to the high rate of interactions  
1494 ( $O(200)$  events per hour) as well as having similar energies to muons from CCQE  
1495  $\nu_\mu$  interactions from beam interactions. The rate of cosmic muons does depend on  
1496 the solar activity cycle [186] but has been neglected in this comparison study. This  
1497 is because the shape of the distributions is most important for the purposes of being  
1498 compared to the detector systematics. The SK-IV and SK-V data samples consist of  
1499 2398.42 and 626.719 hours of data which equates to 686k and 192k events respectively.



**Figure 5.6.:** The variation of the measured dark rate as a function of date, broken down by PMT type. The SK-IV and SK-V periods span September 2008 to May 2018 and January 2019 to July 2020, respectively. The break in measurement in 2018 corresponds to the period of tank repair and refurbishment. Figure adapted from [185].

The predicted charge calculated in the `fitQun` algorithm includes a contribution from the photoelectron emission due to dark noise. Therefore, the increase in the SK-V dark rate needs to be accounted for. In practice, the average dark rate in each SK period is calculated and used as an input in the reconstruction. This is calculated by averaging the dark rate per run for each period separately, using the calibration measurements detailed in subsection 3.1.2. The average dark rate from SK-IV and SK-V were found to be 4.57kHz and 6.30kHz, respectively. The charges associated with the muon and decay electron subevents are illustrated in Figure 5.7. The photoelectron emission from dark noise is more significant for events that have lower energy. This is because this contribution becomes more comparable to the number of photoelectrons emitted from incident photons in lower-energy events. This behaviour is observed

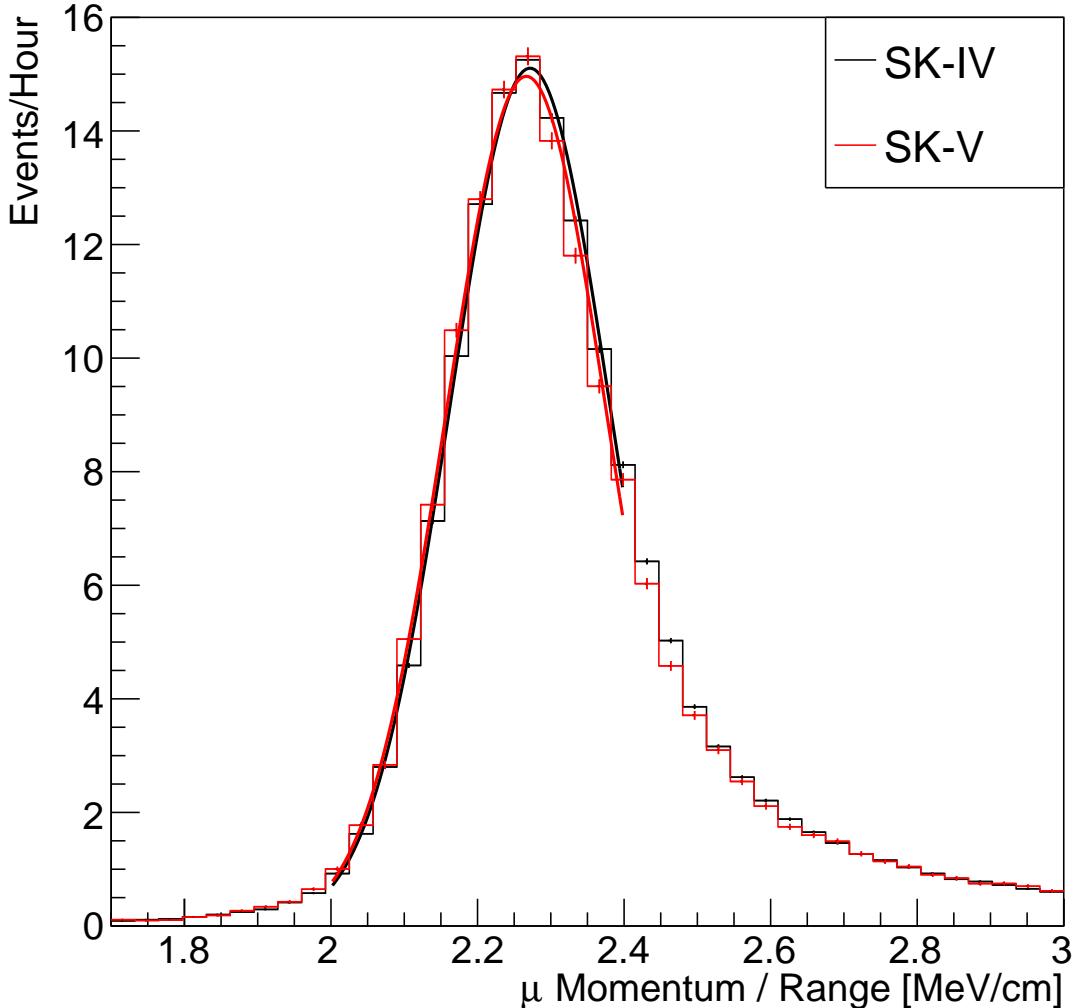
in the data, where the charge deposited by the muon subevent is mostly unaffected by the increase in dark rate, whilst the charge associated with the decay-electron is clearly affected.



**Figure 5.7.:** Comparison of the measured raw charge deposited per event from the stopping muon data samples between SK-IV (Blue) and SK-V (Red), split by the primary muon subevent and the associated decay electron subevent.

The energy scale systematic is estimated from data-to-Monte Carlo differences in the stopping muon sample in [187] and found to be 2.1%. To determine the consistency of SK-IV and SK-V with respect to the energy scale systematic, the muon momentum distribution is compared between the two SK periods. As the total number of Cherenkov photons is integrated across the track length, the reconstructed momentum divided by track length (or range) is compared between SK-IV and SK-V as illustrated in Figure 5.8.

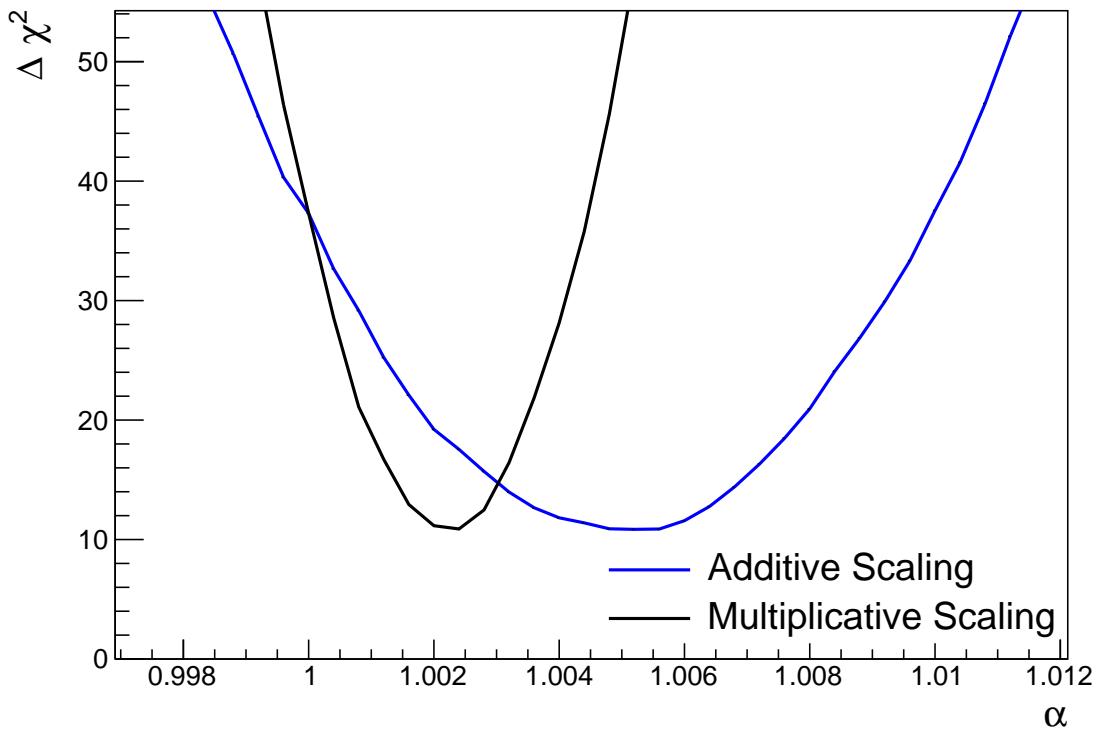
The consistency between these distributions has been computed in two ways. Firstly, a Gaussian is fit to the peak of each distribution separately, whose mean is found to be  $(2.272 \pm 0.003)\text{MeV/cm}$  and  $(2.267 \pm 0.006)\text{MeV/cm}$  for SK-IV and SK-V respectively. The ratio of these is equal to  $1.002 \pm 0.003$ . The means of the



**Figure 5.8.:** The distribution of the reconstructed momentum from the muon ring divided by the distance between the reconstructed muon and decay electron vertices as found in the stopping muon data sets of SK-IV (Black) and SK-IV (Red). Only events with one tagged decay electron are considered. A Gaussian fit is considered in the range [2.0, 2.4] MeV/cm and illustrated as the solid curve.

1525 Gaussian fits are consistent with the expected stopping power of a minimum ionising  
 1526 muon for a target material (water) with  $Z/A \sim 0.5$  [188]. The second consistency  
 1527 check is performed by introducing a nuisance parameter,  $\alpha$ , which modifies the SK-  
 1528 V distribution. The value of  $\alpha$  which minimises the  $\chi^2$  value between the SK-IV  
 1529 and SK-V is determined by scanning across a range of values. This is repeated by  
 1530 applying the nuisance parameter as both a multiplicative factor and an additive

1531 shift. The  $\chi^2$  distributions for different values of  $\alpha$  is illustrated in Figure 5.9. The  
1532 values which minimise the  $\chi^2$  are found to be 0.0052 and 1.0024 for the additive and  
1533 multiplicative implementations, respectively. No evidence of shifts larger than the  
1534 2.1% uncertainty on the energy scale systematic has been found in the reconstructed  
1535 momentum distribution of SK-IV and SK-V.

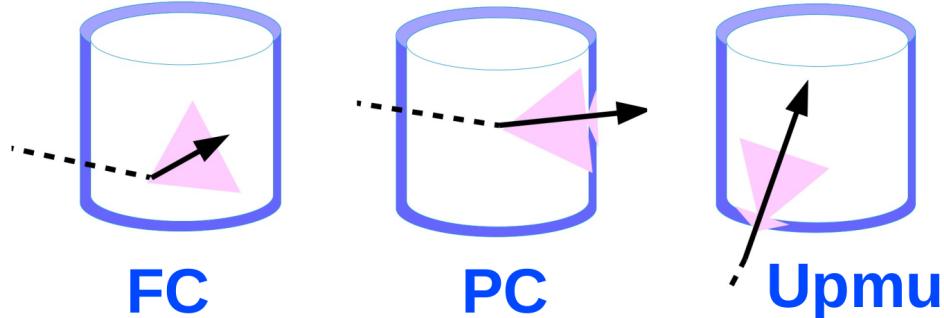


**Figure 5.9.:** The  $\chi^2$  difference between the SK-IV and SK-V reconstructed muon momentum divided by range when the SK-V is modified by the scaling parameter  $\alpha$ . Both additive (Blue) and multiplicative (Black) scaling factors have been considered. In practice, the additive scaling factor actually uses the value of  $(\alpha - 1.0)$  but is illustrated like this so the results can be shown on the same axis range.

### 1536 5.3. Event Reduction at SK

1537 Atmospheric neutrino events observed in the SK detector are categorised into three  
1538 different types of samples: fully contained (FC), partially contained (PC) and up-going

1539 muon (Up- $\mu$ ), using PMT hit signatures in the inner and outer detector (ID and OD,  
 1540 respectively). To identify FC neutrino events, it is required that the neutrino interacts  
 1541 inside the fiducial volume of the ID and that no significant OD activity is observed. For  
 1542 this analysis, an event is defined to be in the fiducial volume provided the event vertex  
 1543 is at least 0.5m away from the ID walls. PC events have the same ID requirements but  
 1544 can have a larger signal present inside the OD. Typically, only high energy muons from  
 1545  $\nu_\mu$  interactions can penetrate the ID wall. The Up- $\mu$  sample contains events where muons  
 1546 are created from neutrino interactions in the OD water or rock below the tank. They  
 1547 then propagate upwards through the detector. Downward-going muons generated  
 1548 from neutrino interactions above the tank are neglected because of the difficulty in  
 1549 separating their signature from the cosmic muon shower background. The sample  
 1550 categories are visually depicted in Figure 5.10.



**Figure 5.10.:** A depiction of the topology patterns for fully-contained (FC), partially-contained (PC) and up-going muon (Up- $\mu$ ) samples included in this analysis.

1551 Based on the event characteristics, as defined by the `fitQun` event reconstruction  
 1552 software, the FC events are categorised by

- 1553 • **Visible Energy:** equal to the sum of the reconstructed kinetic energy of particles  
 1554 above the Cerenkov threshold for all rings present in the event. The purpose is to  
 1555 separate events into sub-GeV and multi-GeV categories.
- 1556 • **Number of observed Cerenkov rings.** The purpose is to separate single-ring and  
 1557 multi-ring events, where single-ring events predominantly consist of quasi-elastic

1558 interactions and multi-ring events are typically resonant pion production or deep  
1559 inelastic scattering events.

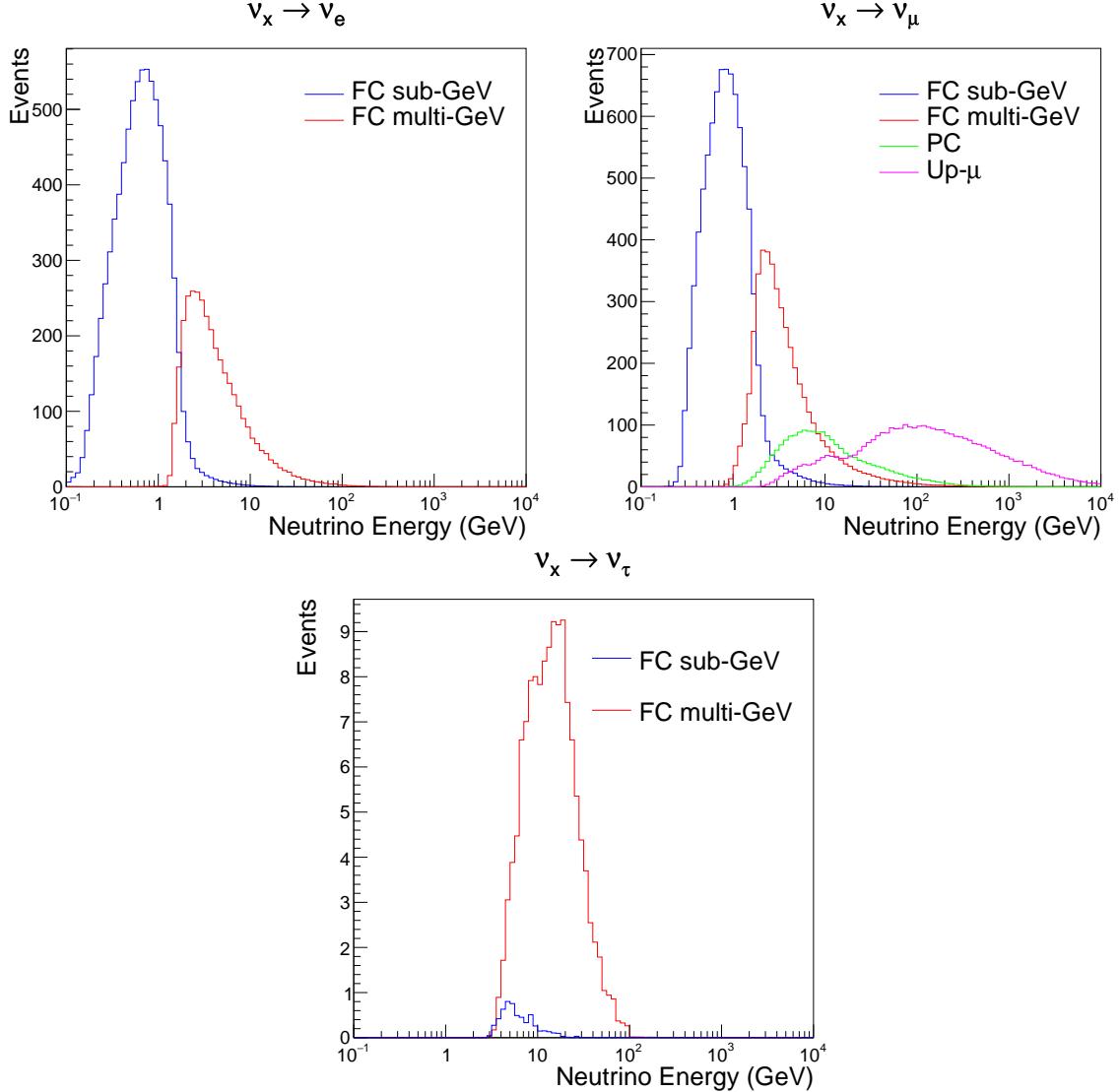
- 1560 • **Particle identification parameter of the most energetic ring:** A value deter-  
1561 mined from the maximum likelihood value based on `fitQun`'s electron, muon, or  
1562 pion hypothesis. The purpose is to separate electron-like and muon-like events.  
1563 • **Number of decay electrons:** The purpose is to separate quasi-elastic events  
1564 (which have one decay electron emitted from the muon decay) and resonant pion  
1565 production events (which have two decay electrons emitted from the muon and  
1566 pion).

1567 The PC and Up- $\mu$  categories are broken down into “through-going” and “stopping”

1568 samples depending on whether the muon leaves the detector. This is because the PC  
1569 stopping events deposit the entire energy of the interaction into the detector, resulting  
1570 in better reconstruction. The energy of events that exit the detector has to be estimated,  
1571 with typically worse resolution, which introduces much larger systematic uncertainties.

1572 Through-going Up- $\mu$  samples are further broken down by whether any hadronic  
1573 showering was observed in the event which typically indicates DIS interactions. The  
1574 expected neutrino energy for the different categories is given in Figure 5.11. FC sub-  
1575 GeV and multi-GeV events peak around 0.7GeV and 3GeV respectively, with slightly  
1576 different peak energies for  $\nu_e$  and  $n\nu_\mu$  oscillation channels. PC and Up- $\mu$  are almost  
1577 entirely comprised of  $\nu_\mu$  events and peak around 7GeV and 100GeV, respectively.

1578 In normal data-taking operations, the SK detector observes many background  
1579 events alongside the beam and atmospheric neutrino signal events of physics interest  
1580 for this thesis. Cosmic ray muons and flasher events, which are the spontaneous  
1581 discharge of a given PMT, contribute the largest amount of background events in the  
1582 energy range relevant to this thesis. Lower energy analyses like DSNB searches are  
1583 also subject to radioactive backgrounds [189]. Therefore the data recorded is reduced



**Figure 5.11.:** The predicted neutrino flux of the fully contained (FC) sub-GeV and multi-GeV, partially contained (PC), and upward-going muon (Up- $\mu$ ) events. The prediction is broken down by the  $\nu_x \rightarrow \nu_e$  prediction (top left),  $\nu_x \rightarrow \nu_\mu$  prediction (top right) and  $\nu_x \rightarrow \nu_\tau$  prediction (bottom). Asimov A oscillation parameters are assumed (given in Table 2.2).

1584 with the aim of removing these background events. The reduction process is detailed  
 1585 in [47, 89] and briefly summarised below.

1586 The first two steps in the FC reconstruction remove the majority of cosmic ray  
 1587 muons by requiring a significant amount of ID activity compared to that measured in  
 1588 the OD. Events that pass this cut are typically very high momentum muons or events  
 1589 that leave very little activity in the OD. Consequently, a third reduction step is then

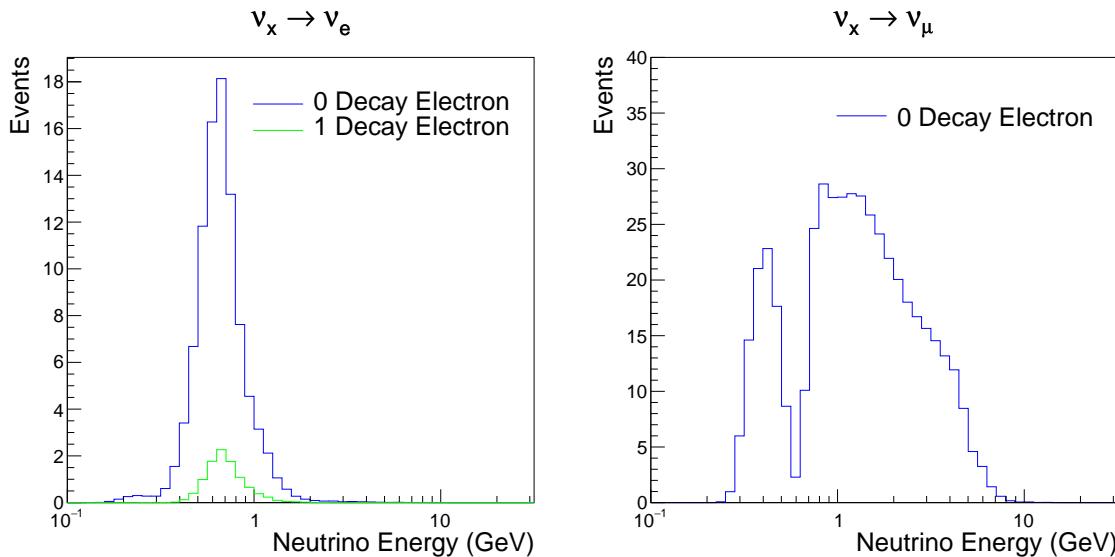
1590 applied to select cosmic-ray muons that pass the initial reduction step. A purpose-built  
1591 cosmic muon fitter is used to determine the entrance (or exit) position of the muon and  
1592 a cut is applied to OD activity contained within 8m of this position. Flasher events are  
1593 removed in the fourth reduction step which is based on the close proximity of PMT  
1594 hits surrounding the PMT producing the flash. Events that pass all these reduction  
1595 steps are reconstructed with the APFit algorithm. The fifth step of the reduction uses  
1596 information from the more precise fitter to repeat the previous two steps with tighter  
1597 cuts. Muons below the Cherenkov threshold can not generate optical photons in the  
1598 ID but the associated decay electron can due to its lower mass. These are the types of  
1599 events targeted in the fifth reduction step. The final cuts require the event vertex to be  
1600 within the fiducial volume (0.5m from the wall although the nominal distance is 2.0m),  
1601 visible energy  $E_{vis} > 30\text{MeV}$  and fewer than 16 hits within the higher energy OD  
1602 cluster. The culmination of the fully contained reduction results in 8.09 events/day in  
1603 the nominal fiducial volume [190]. The uncertainty in the reconstruction is calculated  
1604 by comparing Monte Carlo prediction to data. The largest discrepancy is found to be  
1605 1.3% in the fourth reduction step.

1606 The PC and Up- $\mu$  events are processed through their own reduction processes  
1607 detailed in [47]. Both of these samples are reconstructed with the APFit algorithm  
1608 rather than fitQun. This is because the efficiency of reconstructing events that leave  
1609 the detector has not been sufficiently studied for reliable systematic uncertainties wth  
1610 fitQun. The PC and Up- $\mu$  samples acquire events at approximately 0.66 and 1.44  
1611 events/day.

1612 Events due to beam neutrinos undergo the same reduction steps as FC events and  
1613 are then subject to further cuts [191]. The GPS system which links the timing between  
1614 the beam facility and SK needs to be operating correctly and there should be no activity

<sub>1615</sub> within the detector in the previous  $100\mu\text{s}$  before the trigger. The events then need to  
<sub>1616</sub> triggered between  $-2\mu\text{s}$  and  $10\mu\text{s}$  of the expected spill timing.

<sub>1617</sub> The beam neutrino samples are not split by visible energy since their energy range  
<sub>1618</sub> is smaller than the atmospheric neutrino events. Following the T2K analysis in [80],  
<sub>1619</sub> only single-ring beam neutrino events are considered. Similar to atmospheric event  
<sub>1620</sub> selection, the number of decay electrons is used as a proxy for distinguishing CCQE  
<sub>1621</sub> and CCRES events. The expected neutrino energy, broken down by number of decay  
<sub>1622</sub> electrons, is given in Figure 5.12.



**Figure 5.12.:** The predicted flux of beam neutrinos, as a function of neutrino energy. The predictions are broken down by the number of decay electrons associated with the particular events. Asimov A oscillation parameters are assumed (given in Table 2.2).

<sub>1623</sub> **Chapter 6**

<sub>1624</sub> **Sample Selections and Systematics**

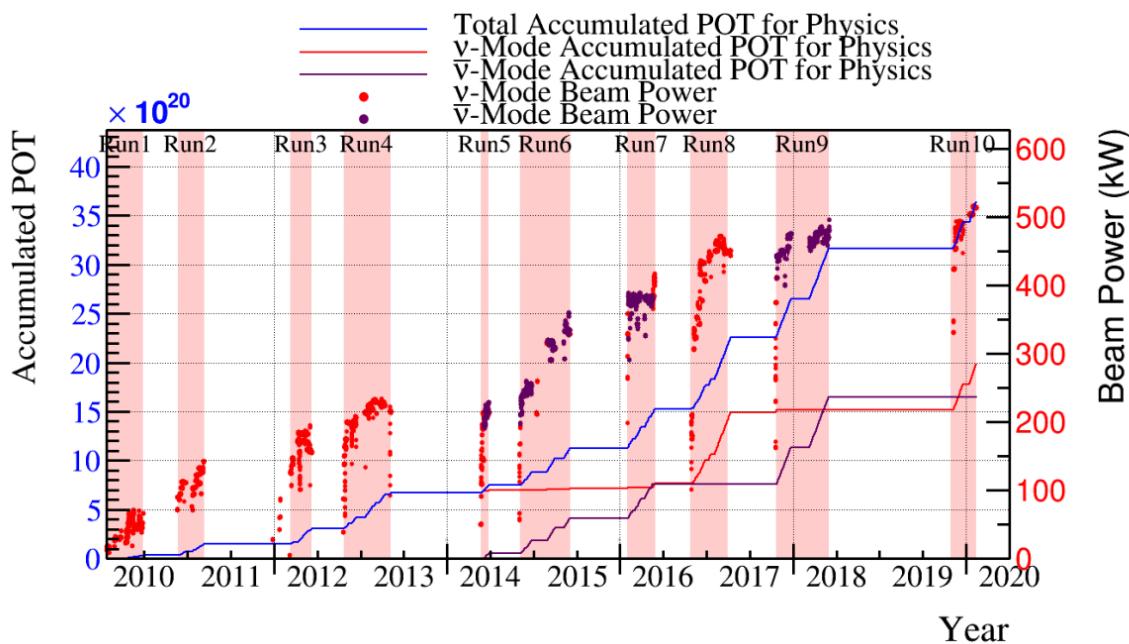
<sub>1625</sub> The oscillation analysis presented within this thesis is built upon a simultaneous  
<sub>1626</sub> fit to atmospheric data at SK, neutrino beam data in the near detector, and beam  
<sub>1627</sub> data measured at SK. The definitions of these samples are documented in section 6.1,  
<sub>1628</sub> section 6.2, and section 6.3, respectively. The data collected and used within this  
<sub>1629</sub> analysis is detailed in Table 6.1. The near and far detector data corresponds to T2K  
<sub>1630</sub> runs 2-9 and runs 1-10, respectively. The accumulated POT and beam power for runs  
<sub>1631</sub> 1 – 10 are illustrated in Figure 6.1.

Data Type	Total
Near Detector FHC	$1.15 \times 10^{21}$ POT
Near Detector RHC	$8.34 \times 10^{20}$ POT
Far Detector FHC	$1.97 \times 10^{21}$ POT
Far Detector RHC	$1.63 \times 10^{21}$ POT
Atmospheric SK-IV	3244.4 days

**Table 6.1.:** The amount of data collected in each detector used within this analysis. The data collected at the near and far detector, for both neutrino beam (FHC) and antineutrino beam (RHC), is measured as the number of protons on target (POT).

<sub>1632</sub> The difference in POT recorded at the near and far detector is due to the difference  
<sub>1633</sub> in downtime. The SK detector is very stable with almost 100% of data recorded during  
<sub>1634</sub> beam operation. Due to various technical and operational issues, the downtime of  
<sub>1635</sub> the near detector is significantly higher due to its more complex design and operating  
<sub>1636</sub> requirements.

1637 The systematic parameters invoked within the flux, detector, and interaction models  
 1638 used within this analysis are documented in section 6.4. The standard configuration of  
 1639 the joint beam and atmospheric data fit utilises far detector systematics provided in  
 1640 the official inputs from the two experiments. Additionally, a correlated detector model  
 1641 which fits the parameters used in sample selections to data has been developed and  
 1642 documented in subsection 6.4.5.



**Figure 6.1.:** The accumulated beam data, measured as the number of protons on target (POT). The total data (blue) is given which comprises of the neutrino beam (red) and antineutrino (purple) components. The beam power for neutrino and antineutrino beams is given as the markers using the same colour scheme. The timescale runs from Run 1 which started in January 2010 until Run 10 which ended in February 2020. The ratio of accumulated data in neutrino and antineutrino beam is 54.7% : 45.3%.

## <sup>1643</sup> 6.1. Atmospheric Samples

<sup>1644</sup> The atmospheric event selection follows the official SK-IV analysis presented in [89]  
<sup>1645</sup> and is documented below. The Monte Carlo prediction used within this analysis  
<sup>1646</sup> corresponds to 500 years worth of neutrino events, which is scaled down to match the  
<sup>1647</sup> SK-IV livetime of 3244.4 days.

<sup>1648</sup> The fully contained (FC), partially contained (PC), and upward going muon events  
<sup>1649</sup> ( $\text{up-}\mu$ ) which pass the reduction cuts discussed in section 5.3 are further broken down  
<sup>1650</sup> into different samples based on reconstruction information. This section details the  
<sup>1651</sup> samples used within this oscillation analysis, alongside the chosen binning,~~used~~  
<sup>1652</sup> [within the fit.](#)

<sup>1653</sup> FC events are first separated by the visible energy deposited within the detector.  
<sup>1654</sup> This is calculated as the sum of the reconstructed kinetic energy above the Cherenkov  
<sup>1655</sup> threshold for all rings present in the event. Events are separated by whether they were  
<sup>1656</sup> above or below  $E_{\text{vis}} = 1.33\text{GeV}$ . This separates “subGeV” and “multiGeV” events.  
<sup>1657</sup> Typically, lower energy events consist of charged current quasi-elastic (CCQE) inter-  
<sup>1658</sup> actions which are better understood and simpler to reconstruct resulting in smaller  
<sup>1659</sup> systematic uncertainties. Events are further separated by the number of rings as-  
<sup>1660</sup> sociated with the event due to similar reasoning. As the oscillation probability is  
<sup>1661</sup> dependent upon the flavour of neutrino, electron and muon events are separated  
<sup>1662</sup> using a similar likelihood method to that discussed in section 5.2. To reduce computa-  
<sup>1663</sup> tional resources required for the reconstruction, only electron and pion hypotheses are  
<sup>1664</sup> considered so this separation cut depends on the ratio of the electron to pion likeli-  
<sup>1665</sup> hoods,  $\log(L_e/L_\pi)$ . Finally, the number of decay electrons is used to classify events.  
<sup>1666</sup> Charged current resonant pion production (CCRES) interactions generate a final-state  
<sup>1667</sup> pion. This can decay, mostly likely through a muon, into a decay electron. Therefore

<sub>1668</sub> any electron-like event with one decay electron or muon-like event with two decay  
<sub>1669</sub> electrons was most likely produced by a CCRES interaction. Consequently, the number  
<sub>1670</sub> of decay electrons can be used to distinguish CCQE and CCRES interaction modes.  
<sub>1671</sub> Ultimately, FC subGeV events are separated into the samples listed in Table 6.2.

Sample Name	Description
SubGeV- <i>e</i> like-0dcy	Single ring <i>e</i> -like events with zero decay electrons
SubGeV- <i>e</i> like-1dcy	Single ring <i>e</i> -like events with one or more decay electrons
SubGeV- <i>μ</i> like-0dcy	Single ring <i>μ</i> -like events with zero decay electrons
SubGeV- <i>μ</i> like-1dcy	Single ring <i>μ</i> -like events with one decay electrons
SubGeV- <i>μ</i> like-2dcy	Single ring <i>μ</i> -like events with two or more decay electrons
SubGeV- <i>π</i> 0like	Two <i>e</i> -like ring events with zero decay electrons and reconstructed $\pi^0$ mass $85 \leq m_{\pi^0} < 215$ MeV

**Table 6.2.:** The fully contained subGeV samples, defined as events with visible energy  $E_{vis} < 1.33$  GeV, used within this oscillation analysis.

<sub>1672</sub> In addition to the cuts discussed above, multiGeV samples also have additional  
<sub>1673</sub> cuts to separate samples which target neutrino and antineutrino **separation events**.  
<sub>1674</sub> As discussed in section 7.1, the matter resonance only occurs for neutrinos in normal  
<sub>1675</sub> hierarchy and antineutrinos in an inverted mass hierarchy. Therefore, having flavour-  
<sub>1676</sub> enriched samples aids in the determination of the mass hierarchy. For a CCRES  
<sub>1677</sub> interaction,

$$\begin{aligned}
 \bar{\nu}_e + N &\rightarrow e^+ + N' + \pi^-, \\
 \nu_e + N &\rightarrow e^- + N' + \pi^+ \\
 &\quad \downarrow \mu^+ + \nu_\mu \\
 &\quad \downarrow e^+ + \nu_e + \bar{\nu}_\mu.
 \end{aligned} \tag{6.1}$$

1678     The  $\pi^-$  emitted from a  $\bar{\nu}_e$  interaction is more likely to be absorbed within the  
1679   oxygen nucleus ~~compared to than~~ the  $\pi^+$  from  $\nu_e$  interactions [192]. These pions  
1680   then decay, mostly through muons, to electrons. Therefore the number of tagged  
1681   decay electrons associated with an event gives an indication of whether the interaction  
1682   was due to a neutrino or antineutrino: zero for  $\bar{\nu}_e$  events, and one for  $\nu_e$  events. The  
1683   ability to separate neutrino from antineutrino events is illustrated in Table 6.4, where  
1684   the MultiGeV-*elike-nue* has 78% purity of CC neutrino interactions with only 7%  
1685   antineutrino background ~~in that sample and the rest of the sample comprising of~~  
1686   ~~neutral current backgrounds, the rest consisting of NC backgrounds.~~

1687     This relatively simple The number of decay electrons discriminator works rea-  
1688   sonably well for single-ring events. However, this is not the case for multi-ring events.  
1689   A multiGeV multiring ~~electron-like separation~~ (MME) likelihood cut ~~which specifi-~~  
1690   ~~cally targets multiGeV multiRing electron-like events~~ was introduced in [193, 194].  
1691   This is a two-stage likelihood selection cut. Four observables are used in the first  
1692   likelihood cut to distinguish  $CC\nu_e$  and  $CC\bar{\nu}_e$  events from background:

- 1693     • The number of decay electrons
- 1694     • The maximum distance between the vertex of the neutrino and the decay electrons
- 1695     • The energy deposited by the highest energy ring
- 1696     • The particle identification of that highest energy ring

1697     Below Paragraph Re-written

1698     Background events consist of  $CC\nu_\mu$  and NC interactions. Typically, the majority of  
1699   the energy in these background events is carried by the hadronic system. Additionally,  
1700   muons tend to travel further than the pions from  $CC\nu_e$  before decaying. Thus, the  
1701   parameters used within the likelihood cut target these typical background interaction  
1702   kinematics.

Sample Name	Description
MultiGeV-elike-nue	Single ring $e$ -like events with zero decay electrons
MultiGeV-elike-nuebar	Single ring $e$ -like events with one or more decay electrons
MultiGeV-mulike	Single ring $\mu$ -like events
MultiRing-elike-nue	Two or more ring events with leading energy $e$ -like ring and passed both MME and $\nu/\bar{\nu}$ separation cuts
MultiRing-elike-nuebar	Two or more ring events with leading energy $e$ -like ring and passed MME and failed $\nu/\bar{\nu}$ separation cuts
MultiRing-mulike	Two or more ring events with leading energy $\mu$ -like ring and only requires $E_{vis} > 0.6\text{GeV}$ DB: Why is this not 1.33GeV? N
MultiRing-Other1	Two or more ring events with leading energy $e$ -like ring and failed the MME likelihood cut

**Table 6.3.:** The fully contained multiGeV samples used within this oscillation analysis. Both the sample name and description are given.

1703     Neutrino and antineutrino events are then separated by a second likelihood method  
 1704     ( $\nu/\bar{\nu}$  separation) detailed in [52]. This uses the number of decay electrons, the number  
 1705     of reconstructed rings, and the event's transverse momentum. The last two parameters  
 1706     are used because higher-energy samples tend to have more pions produced above  
 1707     the Cherenkov threshold which results in more rings compared to an antineutrino  
 1708     interaction. Furthermore, the angular distribution also tends to be more forward  
 1709     peaked in antineutrino interactions as compared to neutrino interactions [89]. These  
 1710     FC multiGeV sample definitions are detailed in Table 6.3.

1711     **Below Paragraph Re-worded**

1712     The PC and up- $\mu$  **events samples** are split by the amount of energy deposited  
 1713     within the outer detector, into “stopping” and “through-going” samples. If an event  
 1714     leaves the detector, the energy it takes with it has to be estimated which increases  
 1715     the systematic uncertainty compared to events entirely contained within the inner  
 1716     detector. This estimation is particularly poor at high energies, thus the up- $\mu$  through-

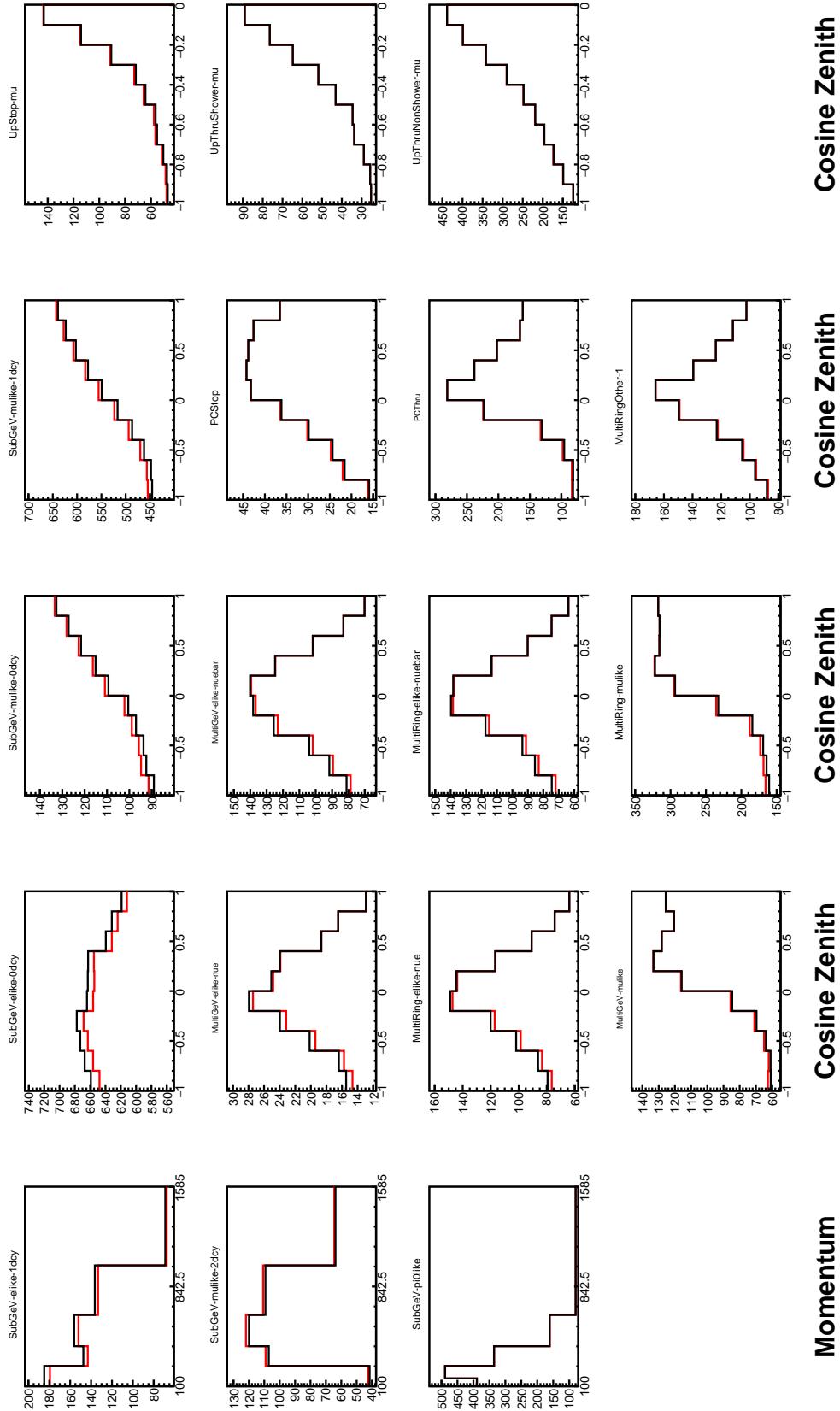
<sup>1717</sup> going events are not binned in reconstructed momentum. The through-going up- $\mu$   
<sup>1718</sup> are further separated by the presence of any electromagnetic showering in the event,  
<sup>1719</sup> as the assumption of non-showering muon does not give reliable reconstruction for  
<sup>1720</sup> these types of events [47]. In total, 13 FC, 2 PC, and 3 up- $\mu$  atmospheric samples are  
<sup>1721</sup> included within this analysis.

Sample	CC $\nu_e$	CC $\bar{\nu}_e$	CC( $\nu_\mu + \bar{\nu}_\mu$ )	CC( $\nu_\tau + \bar{\nu}_\tau$ )	NC
SubGeV- <i>elike</i> -0dcy	72.17	23.3	0.724	0.033	3.77
SubGeV- <i>elike</i> -1dcy	86.81	1.773	7.002	0.062	4.351
SubGeV- <i>mulike</i> -0dcy	1.003	0.380	90.07	0.036	8.511
SubGeV- <i>mulike</i> -1dcy	0.023	0.	98.46	0.029	1.484
SubGeV- <i>mulike</i> -2dcy	0.012	0.	99.25	0.030	0.711
SubGeV- <i>pi0like</i>	6.923	2.368	0.928	0.011	89.77
MultiGeV- <i>elike</i> -nue	78.18	7.041	3.439	1.886	9.451
MultiGeV- <i>elike</i> -nuebar	56.68	37.81	0.174	0.614	4.718
MultiGeV- <i>mulike</i>	0.024	0.005	99.67	0.245	0.058
MultiRing- <i>elike</i> -nue	59.32	12.39	4.906	3.385	20
MultiRing- <i>elike</i> -nuebar	52.39	31.03	1.854	1.585	13.14
MultiRing- <i>mulike</i>	0.673	0.080	97.33	0.342	1.578
MultiRingOther-1	27.98	2.366	34.93	4.946	29.78
PCStop	8.216	3.118	84.45	0.	4.214
PCThrus	0.564	0.207	98.65	0.	0.576
UpStop-mu	0.829	0.370	98.51	0.	0.289
UpThruNonShower-mu	0.206	0.073	99.62	0.	0.103
UpThruShower-mu	0.128	0.054	99.69	0.	0.132

**Table 6.4.:** The purity of each atmospheric sample used within this analysis, broken down by charged current (CC) and neutral current (NC) interactions and which neutrino flavour interacted within the detector. Asimov A oscillation parameter sets are assumed (given in Table 2.2). Electron neutrino and antineutrino events are separated to illustrate the ability of the separation likelihood cuts used within the multiGeV and multiring sample selections.

<sup>1722</sup> The atmospheric samples are binned in direct observables: reconstructed lepton  
<sup>1723</sup> momentum and direction, as given by Table A.1. The distribution of the reconstructed  
<sup>1724</sup> lepton momentum (for samples that only have one bin in reconstructed zenith angle)

<sub>1725</sub> and reconstructed direction for each atmospheric sample used within this analysis is  
<sub>1726</sub> illustrated in Figure 6.2. The by-mode breakdown of each of the atmospheric samples  
<sub>1727</sub> is given in Appendix A.



**Figure 6.2:** Comparison of the SK-IV atmospheric samples between predictions made with the CP-violating Asimov A (Black) and CP-conserving Asimov B (Red) oscillation parameter sets (given in Table 2.2). The subGeV samples CCRES and  $\pi^0$ -like samples are given in their reconstructed lepton momentum. All other samples are presented in their reconstructed zenith angle projection.

## 1728 6.2. Near Detector Beam Samples

1729 The near detector sample selections are documented in detail within [195] and sum-  
1730 marised below. Samples are selected based upon the which Fine Grained Detector  
1731 (FGD) that the vertex is reconstructed in as well as the operating mode of the beam:  
1732 FHC or RHC. ~~For additional constraints on model parameters, wrong-sign neutrino~~  
1733 ~~samples are also considered when the beam is operating in RHC mode. Wrong-~~  
1734 ~~sign neutrino samples are considered in the RHC mode in order to add additional~~  
1735 ~~constraints on model parameters.~~ Samples from the wrong-sign component of the  
1736 FHC beam mode are not included as they are statistically insignificant compared to  
1737 those samples already listed.

1738 For additional constraints on model parameters, wrong-sign neutrino samples are  
1739 also considered when the beam is operating in RHC mode.

1740 The reconstruction algorithm uses a clustering algorithm to group hits within the  
1741 TPC. It then adds information from the upstream FGD to form a track which passes  
1742 through both sub-detectors. In FHC(RHC), the highest momentum negative(positive)  
1743 curvature track is defined as the muon candidate. Before being assigned a sample,  
1744 these candidate muon events must ~~undergo pass~~ CC-inclusive cuts, as defined in [196]:

- 1745 • Event Timing: The DAQ must be operational and the event must occur within  
1746 the expected beam time window consistent with the beam spill
- 1747 • TPC Requirement: The muon-candidate track path must intercept one or more  
1748 TPCs
- 1749 • Fiducial volume: The event must originate from within the fiducial volume .  
1750 ~~The fiducial volumes are defined as a region within each sub-detector defined~~  
1751 ~~in [197].~~

1752 • Upstream Background: Remove events that have muon tracks that originate  
1753 upstream of the FGDs by requiring no high-momentum tracks within 150mm  
1754 upstream of the candidate vertex. Additionally, events that occur within the  
1755 downstream FGD are vetoed if a secondary track starts within the upstream FGD

1756 • Broken track removal: All candidates where the muon candidate is broken in two  
1757 are removed

1758 • Muon PID: Measurements of  $dE/dx$  in a TPC are used to distinguish muon-like  
1759 events **from electron-like or proton-like**, using a likelihood cut

1760 In addition to these cuts, RHC neutrino events also have to undergo the following  
1761 cuts to aid in the separation of neutrino and antineutrino [198]:

1762 • TPC Requirement: The track path must intercept TPC2

1763 • Positive Track: The highest momentum track must have positive reconstructed  
1764 charge

1765 • TPC1 Veto: Remove any events originating upstream of TPC1

1766 Once all CC-inclusive events have been determined, they are further **separated**  
1767 ~~into sub-samples that target the constraints on interaction modes most relevant at~~  
1768 ~~the far-detector. They are~~ split by pion multiplicity: CC0 $\pi$ , CC1 $\pi$ , and CCOther.

1769 These target specific interaction modes CCQE, CCRES, and other CC background  
1770 interactions, respectively. Pions in the TPCs and FGDs are selected by requiring a  
1771 second track to be observed, which is separate from the muon track and is in the same  
1772 beam spill window and sub-detector. If the pion originated within a FGD, it must also  
1773 pass through the sequential downstream TPC (TPC2 for FGD1, TPC3 for FGD2).

1774 CC0 $\pi$ , CC1 $\pi$ , and CCOther samples are defined with the following cuts:

1775 DB: Understand pion cuts at ND

1776 •  $\nu_\mu$ **CC0 $\pi$  Selection:** No electrons in TPC and no charged pions or decay electrons

1777 within the TPC or FGD

1778 •  $\nu_\mu$ **CC1 $\pi$  Selection:** Exactly one charged pion in either the TPC or FGD, where

1779 the number of charged pions in the FGD is equal to the number of decay electrons

1780 •  $\nu_\mu$ **CCOther Selection:** All events which are not classified into the above two

1781 selections.

1782 Counting the three selections for each FGD in FHC and RHC running, including

1783 the wrong-sign background in RHC, 18 near detector samples are used within this

1784 analysis. These samples are binned in reconstructed lepton momentum (illustrated in

1785 Figure 6.3) and direction with respect to the beam. The binning is chosen such that

1786 each event has at least 20 Monte Carlo events in each bin [197]. This is to ensure that

1787 the bins are coarse enough to ensure the reduction of statistical errors, whilst also

1788 being fine enough to sample the high-resolution peak regions. The exact binning is

1789 detailed in [197].

### 1790 6.3. Far Detector Beam Samples

1791 The beam neutrino events which occur at the SK detector, which pass the reduction

1792 cuts detailed in section 5.3, are separated **depending based** on whether the beam was

1793 operating in FHC or RHC mode. The events are then separated into three samples:

1794 electron-like (1Re), muon-like (1R $\mu$ ), and CC1 $\pi^+$ -like (1Re1de) which are observed as

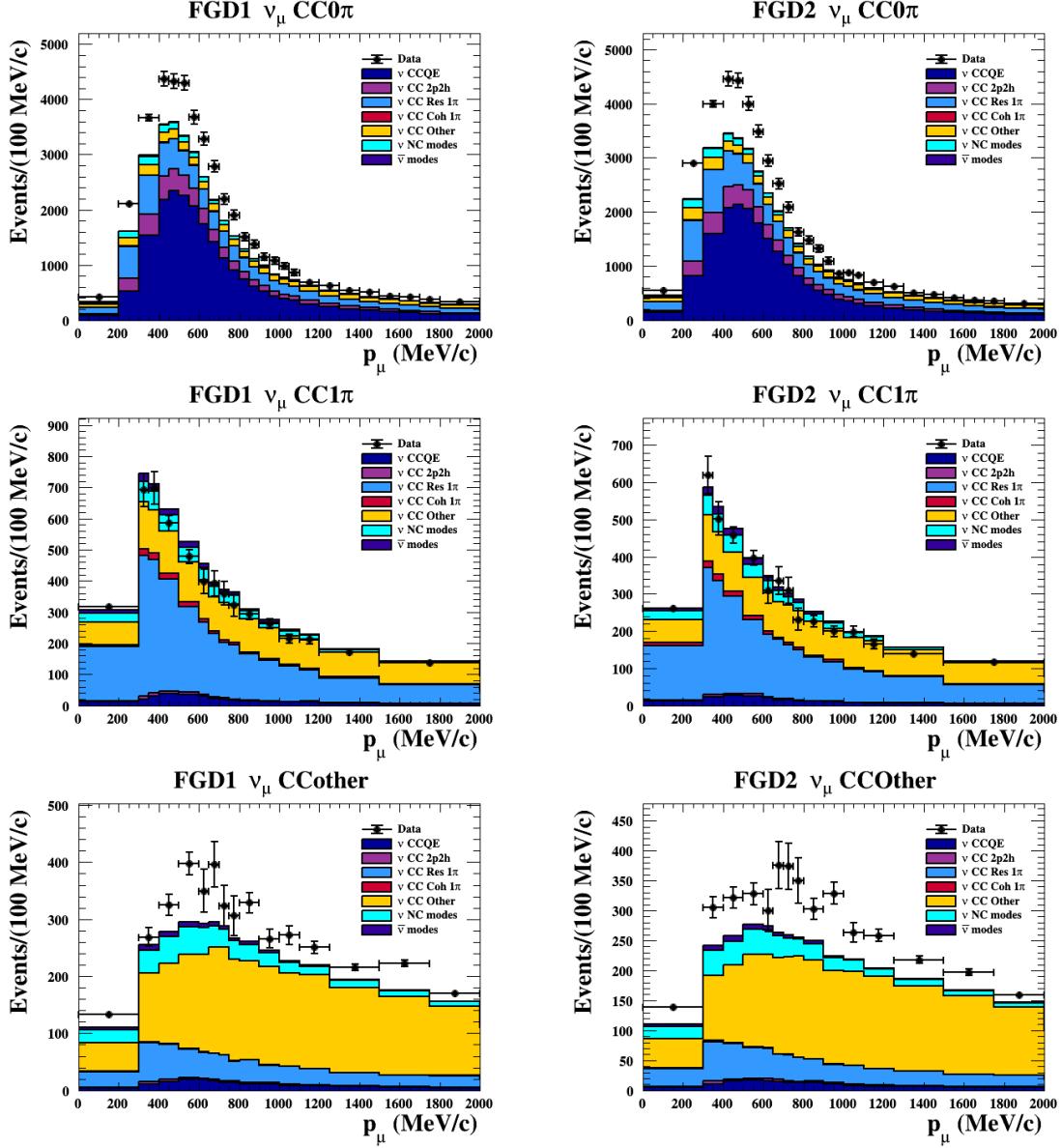
1795 electron-like events with an associated decay electron [185]. As discussed in section 6.1,

1796 positively charged pions emitted from neutrino interactions are more likely to produce

1797 decay electrons than negatively charged pions. Consequently, the CC1 $\pi^+$ -like sample

1798 is only selected when the beam is operating in FHC mode. Therefore, five beam

1799 samples measured at SK are used in this analysis.



**Figure 6.3.:** The nominal Monte Carlo predictions compared to data for the FGD1 and FGD2 samples in neutrino beam mode, broken down into the  $CC\nu_\mu 0\pi$ ,  $CC\nu_\mu 1\pi$  and  $CC\nu_\mu$  Other categories. Figures taken from [195].

1800        The fiducial volume definition for beam samples is slightly different from that  
 1801        used **within for** the atmospheric samples. It uses both the distance to the closest wall  
 1802        (`dWall`) and the distance to the wall along the trajectory of the particle (`toWall`). This  
 1803        allows events that originate close to the wall but are facing into the tank to be included  
 1804        within the analysis, which would have otherwise been removed. These additional  
 1805        events are beneficial for a statistics-limited experiment. The exact cut values for both

1806 dWall and toWall are different for each of the three types of sample and are optimised  
1807 based on T2K sensitivity to  $\delta_{CP}$  [183, 199]. They are:

1808 **1Re event selection** For an event to be classified as a 1Re-like, the event must **follow**  
1809 **satisfy**:

- 1810 • Fully-contained and **within have**  $d_{Wall} > 80\text{cm}$  and  $t_{Wall} > 170\text{cm}$
- 1811 • Total of one ring which is reconstructed as electron-like with reconstructed mo-
- 1812      mentum  $P_e > 100\text{MeV}$
- 1813 • Zero decay electrons are associated with the event
- 1814 • Passes  $\pi^0$  rejection cut discussed in section 5.2

1815 **The zero decay electron cut removes non-CCQE interactions and The zero-de-**  
1816 **cay electron cut specifically targets CCQE interactions. Whereas,** the  $\pi^0$  rejection  
1817 cut is designed to remove neutral current  $\pi^0$  background events which can be easily  
1818 reconstructed as 1Re-like events.

1819 The zero decay electron cut removes non-CCQE interactions and the  $\pi^0$  rejection  
1820 cut is designed to remove neutral current  $\pi^0$  background events which can be easily  
1821 reconstructed as 1Re-like events.

1822 **CC1 $\pi^+$  event selection** This event selection is very similar to that of the 1Re sample.  
1823 The only **difference is differences are** that the dWall and toWall criteria are changed  
1824 to  $> 50\text{cm}$  and  $> 270\text{cm}$ , respectively. **Furthermore, , and** exactly one decay electron  
1825 is required from the  $\pi^+$  decay.

1826 **1R $\mu$  event selection** A 1R $\mu$ -like event is determined by the following cuts:

- 1827 • Fully-contained and **within have**  $d_{Wall} > 50\text{cm}$  and  $t_{Wall} > 250\text{cm}$
- 1828 • Total of one ring which is reconstructed as muon-like with reconstructed momentum  $P_\mu > 200\text{MeV}$
- 1829
- 1830 • Fewer than two decay electrons are associated with the event
- 1831 • Passes  $\pi^+$  rejection cut discussed in section 5.2
- 1832 All of these samples are binned in reconstructed neutrino energy. This is possible
- 1833 under a particular interaction mode assumption, as the direction from the source is
- 1834 known extremely well. **This value is calculated** For the 1Re-like and 1R $\mu$ -like samples
- 1835 **assuming CCQE interactions,**

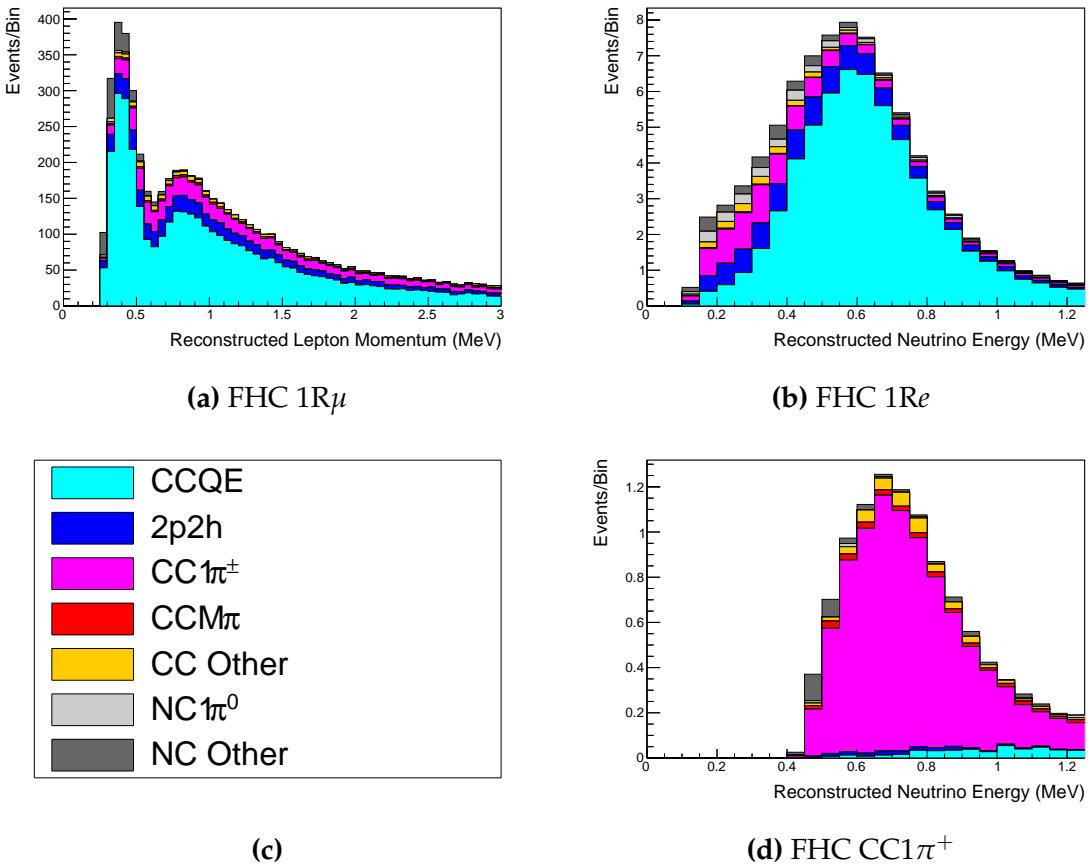
$$E_\nu^{rec} = \frac{(M_N - V_{nuc})E_l - m_l^2/2 + M_N V_{nuc} - V_{nuc}^2/2 + (M_P^2 + M_N^2)/2}{M_N - V_{nuc} - E_l + P_l \cos(\theta_{beam})} \quad (6.2)$$

1836 Where  $M_N$ ,  $M_P$  and  $m_l$  are the masses of the neutron, proton and outgoing lepton,  
 1837 respectively.  $V_{nuc} = 27\text{MeV}$  is the binding energy of the oxygen **nuclei nucleus** [185],  
 1838  $\theta_{beam}$  is the angle between the beam and the direction of the outgoing lepton, and  $E_l$   
 1839 and  $P_l$  are the energy and momentum of that outgoing lepton.

1840 The reconstructed neutrino energy of the CC1 $\pi^+$ -like events **is modified to include**  
 1841 **also accounts for** the delta resonance produced within the interaction,

$$E_\nu^{rec} = \frac{2M_N E_l + M_{\Delta^{++}}^2 - M_N^2 - m_l^2}{2(M_N - E_l + P_l \cos(\theta_{beam}))} \quad (6.3)$$

1842 Where  $M_{\Delta^{++}}$  is the mass of the delta baryon. Binding energy effects are not  
 1843 considered as a two-body process with the delta baryon is assumed. This follows

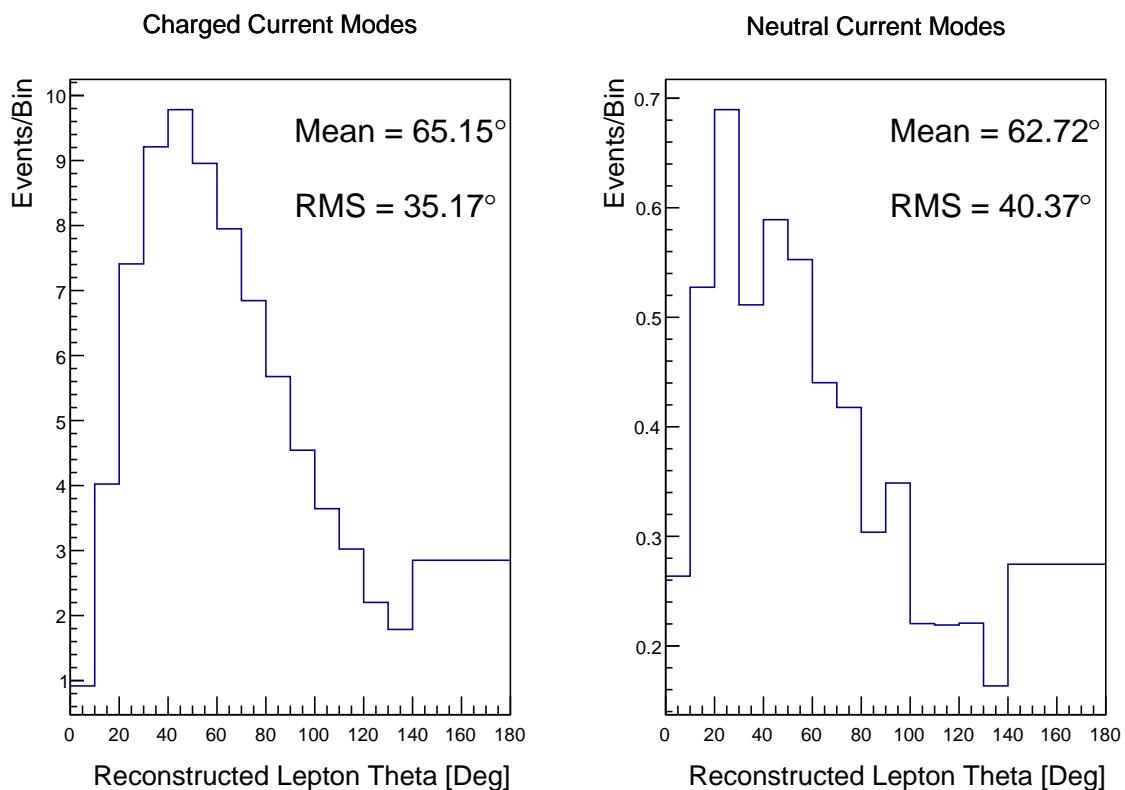


**Figure 6.4.:** The reconstructed neutrino energy, as defined by Equation 6.2 and Equation 6.3, for the  $1R\mu$ -like,  $1Re$ -like and  $CC1\pi^+$ -like samples. Asimov A oscillation parameter sets are assumed (given in Table 2.2). These samples are the FHC mode samples. For ease of viewing, the  $1R\mu$  sample only shows the  $0. \leq E_\nu^{rec} < 3.0$ GeV but the binning extends to 30.0GeV.

1844 the T2K oscillation analysis presented in [80], although recent developments of the  
 1845 interaction model in the latest T2K oscillation analysis do include effects from binding  
 1846 energy in this calculation [200].

1847 The reconstructed neutrino energy for the FHC samples is illustrated in Figure 6.4.  
 1848 As expected, the  $1R\mu$ -like and  $1Re$ -like samples are heavily dominated by CCQE in-  
 1849 teractions, with smaller contributions from 2p2h meson exchange and resonant pion  
 1850 production interactions. The  $CC1\pi^+$ -like sample predominantly consists of charged  
 1851 current resonant pion production interactions. The  $1Re$ -like and  $CC1\pi^+$ -like samples  
 1852 are also binned by the angle between the neutrino beam and the reconstructed lepton

momentum. This is to aid in charged current and neutral current separation, as indicated in Figure 6.5. This is because the neutral current backgrounds are predominantly due to  $\pi^0$ -decays, where the opening angle of the two gammas alongside the different final state kinematics produces a slightly broader angular distribution compared to the **final state particles electron** originating from charged current  $\nu_e$  interactions.



**Figure 6.5.:** The distribution of the angle between the neutrino beam direction and the reconstructed final state lepton, for the FHC 1Re-like sample. The distribution is broken down by neutrino interaction mode into charged current (left) and neutral current (right) components. Asimov A oscillation parameter sets are assumed (given in Table 2.2). The RMS of the charged and neutral current plots are  $35.17^\circ$  and  $40.37^\circ$ , respectively.

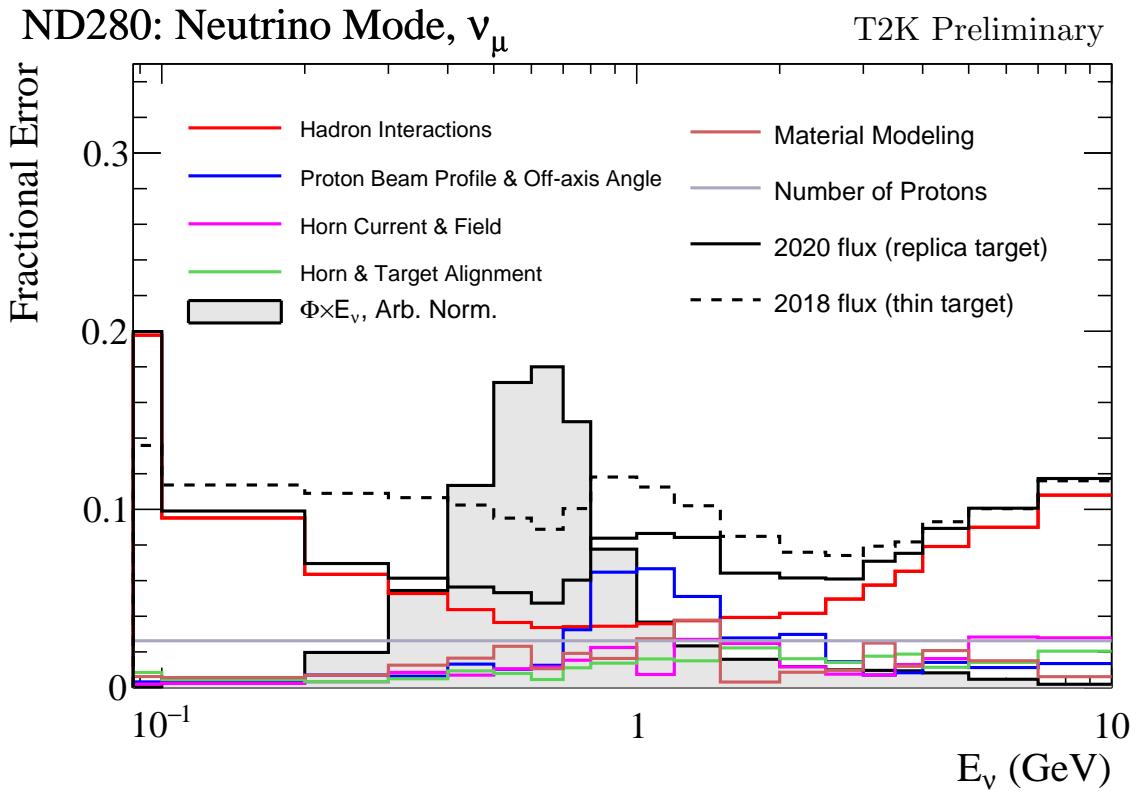
## 1858 6.4. Systematic Uncertainties

1859 The systematic model parameters for this analysis are split into groups, or blocks,  
1860 depending on their purpose. They consist of flux uncertainties, neutrino-matter  
1861 interaction systematics, and detector efficiencies. There are also uncertainties on the  
1862 oscillation parameters which this analysis will not be sensitive to,  $\Delta m_{12}^2$  and  $\sin^2(\theta_{12})$ .  
1863 These uncertainties are taken from the 2018 PDG measurements [81]. As described in  
1864 chapter 4, each model parameter used within this analysis requires a prior uncertainty.  
1865 This is provided via separate covariance matrices for each block. The covariance  
1866 matrices can include prior correlations between parameters within a single block, but  
1867 the separate treatment means prior uncertainties can not be included for parameters  
1868 in different groups. Some parameters in these models have no reasonably motivated  
1869 uncertainties and are assigned flat priors which do not modify the likelihood penalty.  
1870 The flux, neutrino interaction, and detector modeling simulations have already been  
1871 discussed in section 5.1 and section 5.2. The uncertainties invoked within each of these  
1872 models are described below.

### 1873 6.4.1. Beam Flux

1874 The neutrino beam flux systematics is based upon the uncertainty in the modeling of  
1875 the components of the beam. This includes the hadron production model and their re-  
1876 interactions, the shape, intensity, and alignment of the beam with respect to the target,  
1877 and the uniformity of the magnetic field produced by the horn, alongside other effects.  
1878 The uncertainty, as a function of neutrino energy, is illustrated in Figure 6.6 which  
1879 includes a depiction of the total uncertainty as well as the contribution from individual  
1880 components. The uncertainty around the peak of the energy distribution ( $E_\nu \sim 0.6\text{GeV}$ )

<sup>1881</sup> is dominated by the measurements of the beam profile and alignment. Outside of this  
<sup>1882</sup> region, the uncertainties within hadron production dominate the uncertainty.



**Figure 6.6.:** The total uncertainty evaluated on the near detector  $\nu_\mu$  flux prediction constrained by the replica-target data, illustrated as a function of neutrino energy. The solid(dashed) line indicates the uncertainty used within this analysis(the T2K 2018 analysis [201]). The solid histogram indicates the neutrino flux as a function of energy. Figure taken from [202].

<sup>1883</sup> The beam flux uncertainties are described by one hundred parameters. They are  
<sup>1884</sup> split between the ND280 and SK detectors and binned by neutrino flavour:  $\nu_\mu$ ,  $\bar{\nu}_\mu$ ,  $\nu_e$   
<sup>1885</sup> and  $\bar{\nu}_e$ . The response is then broken down as a function of neutrino energy. The bin  
<sup>1886</sup> density in the neutrino energy is the same for the  $\nu_\mu$  in FHC and  $\bar{\nu}_\mu$  in RHC beams,  
<sup>1887</sup> and narrows for neutrino energies close to the oscillation maxima of  $E_\nu = 0.6\text{GeV}$ .  
<sup>1888</sup> This binning is specified in Table 6.5. All of these systematic uncertainties are applied  
<sup>1889</sup> as normalisation parameters with Gaussian priors centered at 1.0 and error specified  
<sup>1890</sup> from a covariance matrix provided by the T2K beam group [202].

Neutrino Flavour	Sign	Neutrino Energy Bin Edges (GeV)
$\mu$	Right	0., 0.4, 0.5, 0.6, 0.7, 1., 1.5, 2.5, 3.5, 5., 7., 30.
$\mu$	Wrong	0., 0.7, 1., 1.5, 2.5, 30.
$e$	Right	0., 0.5, 0.7, 0.8, 1.5, 2.5, 4., 30.
$e$	Wrong	0., 2.5, 30.

**Table 6.5.:** The neutrino energy binning for the different neutrino flavours. “Right” sign indicates neutrinos in the FHC beam and antineutrinos in the RHC beam. “Wrong” sign indicates antineutrinos in the FHC beam and neutrinos in the RHC beam. The binning of the detector response is identical for the FHC and RHC modes as well as at ND280 and SK.

#### 1891 6.4.2. Atmospheric Flux

1892 The atmospheric neutrino flux is modeled by the HKKM model [43]. 16 systematic  
 1893 uncertainties are applied to control the normalisation of each neutrino flavour, energy,  
 1894 and direction. All of the parameters are given Gaussian priors centered at 0 and width  
 1895 equal to one. They are summarised below:

- 1896 • **Absolute Normalisation:** The overall normalisation of each neutrino flavour is  
 1897 controlled by two independent systematic uncertainties, for  $E_\nu < 1\text{GeV}$  and  $E_\nu >$   
 1898  $1\text{GeV}$ , respectively. This is driven mostly by hadronic interaction uncertainties for  
 1899 the production of pions and kaons [43]. The strength of the response is dependent  
 1900 upon the neutrino energy.
- 1901 • **Relative Normalisation:** Uncertainties on the ratio of  $(\nu_\mu + \bar{\nu}_\mu) / (\nu_e + \bar{\nu}_e)$  are  
 1902 controlled by the difference between the HKKM model [43], FLUKA [46] and  
 1903 Bartol models [42]. Three independent parameters are applied in the energy  
 1904 ranges:  $E_\nu < 1\text{GeV}$ ,  $1\text{GeV} < E_\nu < 10\text{GeV}$ , and  $E_\nu > 10\text{GeV}$ .
- 1905 •  **$\nu/\bar{\nu}$  Normalisation:** The uncertainties in the  $\pi^+/\pi^-$  (and kaon equivalent) pro-  
 1906 duction uncertainties in the flux of  $\nu/\bar{\nu}$ . The response is applied using the same  
 1907 methodology as the relative normalisation parameters.

1908 • **Up/Down and Vertical/Horizontal Ratio:** Similar to the above two systematics,  
1909 the difference between the HKKM, FLUKA, and Bartol model predictions, as a  
1910 function of  $\cos(\theta_Z)$ , is used to control the normalisation of events as a function of  
1911 zenith angle.

1912 •  **$K/\pi$  Ratio:** Higher energy neutrinos ( $E_\nu > 10\text{GeV}$ ) become dependent upon  
1913 kaon decay as the dominant source of neutrinos. Measurements of the ratio of  
1914  $K/\pi$  [203] are used to control the systematic uncertainty of the expected ratio of  
1915 pion and kaon production.

1916 • **Solar Activity:** As the 11-year solar cycle can affect the Earth's magnetic field,  
1917 the flux of primary cosmic rays varies across the same period. The uncertainty is  
1918 calculated by taking a  $\pm 1$  year variation, equating to a 10% uncertainty for the  
1919 SK-IV period.

1920 • **Atmospheric Density:** The height of the interaction of the primary cosmic rays is  
1921 dependent upon the atmospheric density. The HKKM assumes the US standard  
1922 1976 [153] profile. This systematic controls the uncertainty in that model.

1923 Updates to the HKKM and Bartol models are underway [158] to use a similar  
1924 tuning technique to that used in the beam flux predictions. After those updates, it may  
1925 be possible to include correlations in the hadron production uncertainty systematics  
1926 for beam and atmospheric flux predictions.

#### 1927 6.4.3. Neutrino Interaction

1928 The neutrino interactions which occur within all the detectors are modeled by NEUT.  
1929 The two independent oscillation analyses, T2K-only [204] and the SK-only [52], have  
1930 developed separate interaction models. To leverage the most sensitivity out of this  
1931 simultaneous beam and atmospheric analysis, a correlated interaction model has been

defined. Where applicable, correlations allow the systematic uncertainties applied to the atmospheric samples to be constrained by measurements of the near detector in the beam experiment. This can lead to stronger sensitivity to oscillation parameters as compared to an uncorrelated model.

The low energy T2K systematic model has a more sophisticated treatment of CCQE, CCMEC, and CCRES uncertainties which is due to the purpose-made cross-section measurements made by the near detector. Furthermore, extensive comparisons of this model have been performed to external data [204]. However, the model is not designed for high-energy atmospheric events, like those illustrated in Figure 5.11. Therefore the high energy systematic model from the SK-only analysis is implemented for the relevant multiGeV, PC, and up- $\mu$  samples. The CCQE systematic parameters invoked within the SK high energy model are actually contained within T2K's CCQE model. Consequently, the more sophisticated CCQE and CCMEC T2K model parameters have been incorporated into the high energy model but are uncorrelated from the low energy counterparts.

The high energy systematic model includes parameters developed from comparisons of Nieves and Rein-Seghal models which affect resonant pion producing interactions, comparisons of the GRV98 and CKMT models which control DIS interactions, and hadron multiplicity measurements which modulate the normalisation of multi-pion producing events. The uncertainty of the  $\nu_\tau$  cross-section is particularly large and is controlled by a 25% normalisation uncertainty. These parameters are applied via normalisation or shape parameters. The former linearly scales the weight of all affected Monte-Carlo events, whereas the latter can increase or decrease a particular event's weight depending on its neutrino energy and mode of interaction. The response of the shape parameters is defined by third-order polynomial splines which return a weight for a particular neutrino energy. To reduce computational resources

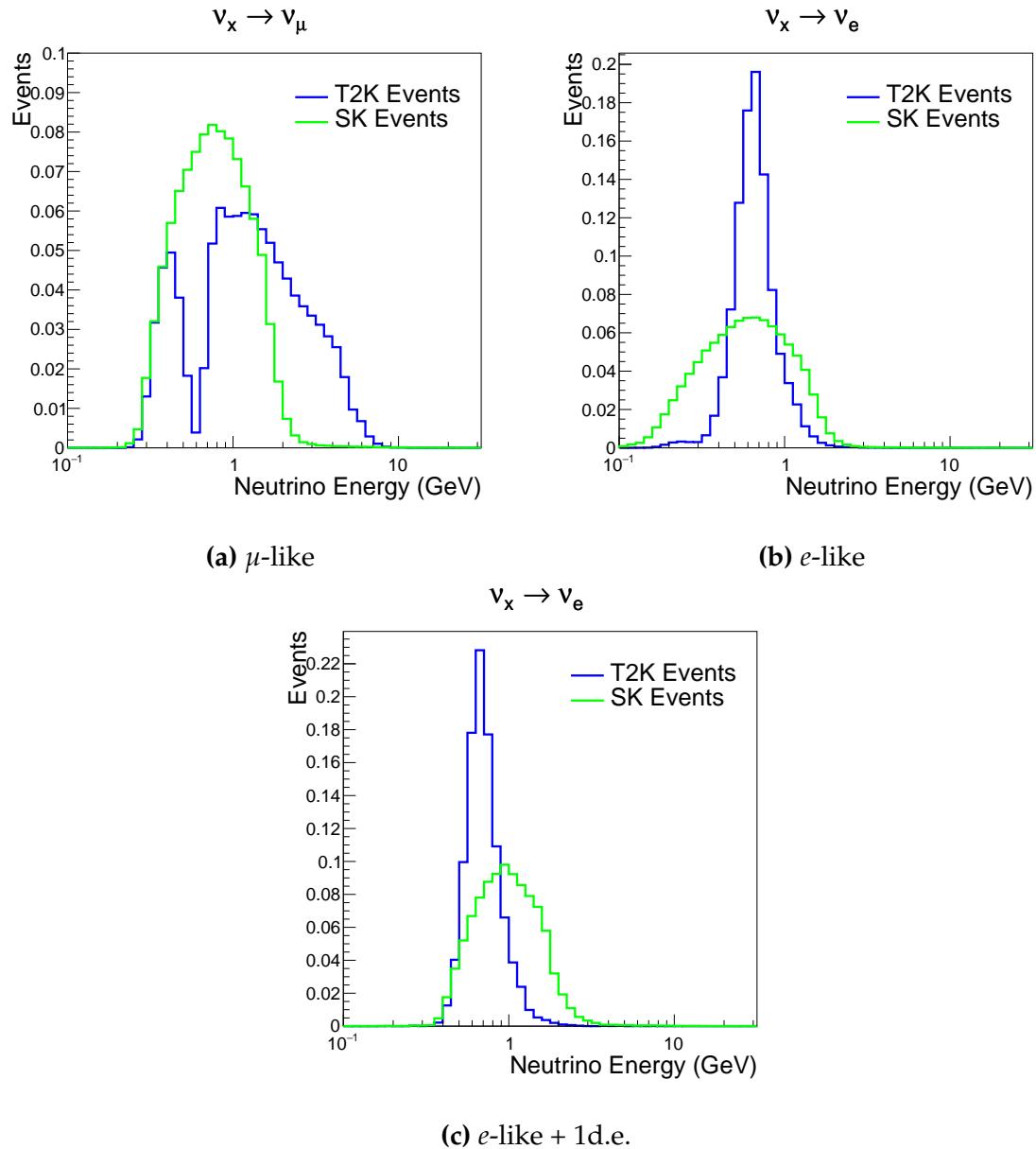
1958 for the far detector fit, the response is binned by neutrino energy and sample binning:  
1959 lepton momentum and cosine zenith binning for atmospheric splined responses and  
1960 reconstructed neutrino energy and direction binning for beam samples. In total, 17  
1961 normalisation and 15 shape parameters are included in the high-energy model within  
1962 this analysis.

1963 Figure 6.7 indicates the predicted neutrino energy distribution for both beam and  
1964 subGeV atmospheric samples. There is clearly significant overlap in neutrino energy  
1965 between the subGeV atmospheric and beam samples, allowing similar kinematics in  
1966 the final state particles. Figure 6.8 illustrates the fractional contribution of the different  
1967 interaction modes per sample.

1968 Comparing beam and atmospheric samples which target CCQE interactions (S.G.  
1969 e-like 0de, S.G.  $\mu$ -like [0,1]de, [FHC,RHC] 1R  $\mu$ -like and [FHC,RHC] 1R e-like  
1970 samples), there is a very similar contribution of CCQE, CC 2p2h, and CC1 $\pi^\pm$  in-  
1971 teractions. The samples which target CC1 $\pi^\pm$  interactions, (S.G. e-like 0de, S.G.  
1972  $\mu$ -like 2de and FHC 1R+1d.e e-like) also consist of very similar mode interactions.

1973 As a consequence of the similarity in energy and mode contributions, correlating  
1974 the systematic model between the beam and subGeV atmospheric samples ensures  
1975 that this analysis attains the largest sensitivity to oscillation parameters while still  
1976 ensuring neutrino interaction systematics are correctly accounted for. Due to its more  
1977 sophisticated CCQE and 2p2h model, the T2K systematic model was chosen as the  
1978 basis of the correlated model.

1979 The T2K systematic model [204] is applied in a similar methodology to the SK model  
1980 parameters. It consists of 19 shape parameters and 24 normalisation parameters. Four  
1981 additional parameters, which model the uncertainty in the binding energy, are applied  
1982 in a way to shift the momentum of the lepton emitted from a nucleus. This controls  
1983 the uncertainty specified on the 27MeV binding energy assumed within Equation 6.2.



**Figure 6.7.:** The predicted neutrino energy distribution for subGeV atmospheric and beam samples. FHC and RHC beam samples are summed together Asimov A oscillation parameters are assumed (given in Table 2.2). Beam and atmospheric samples with similar cuts are compared against one another.

1984 The majority of these parameters are assigned a Gaussian prior uncertainty. Those  
 1985 that have no reasonably motivated uncertainty, or those which have not been fit to  
 1986 external data, are assigned a flat prior which does not affect the penalty term.

1987 There are three particular tunes of the T2K flux and low energy cross section model  
 1988 typically considered. Firstly, the “generated” tune which is the set of dial values with

	CC QE	CC 2p2h	CC $1\pi^\pm$	CC $M\pi$	CC Other	NC $\pi^0$	NC $1\pi^\pm$	NC $M\pi$	NC Coh.	NC Other
FHC 1R+1d.e. e-like	<b>0.04</b>	<b>0.02</b>	<b>0.83</b>	<b>0.03</b>	<b>0.04</b>	<b>0.01</b>	<b>0.01</b>	<b>0.01</b>	<b>0.00</b>	<b>0.01</b>
RHC 1R e-like	<b>0.62</b>	<b>0.12</b>	<b>0.11</b>	<b>0.01</b>	<b>0.02</b>	<b>0.06</b>	<b>0.01</b>	<b>0.01</b>	<b>0.01</b>	<b>0.04</b>
FHC 1R e-like	<b>0.68</b>	<b>0.12</b>	<b>0.10</b>	<b>0.00</b>	<b>0.02</b>	<b>0.04</b>	<b>0.01</b>	<b>0.00</b>	<b>0.00</b>	<b>0.02</b>
RHC 1R $\mu$ -like	<b>0.62</b>	<b>0.13</b>	<b>0.17</b>	<b>0.02</b>	<b>0.03</b>	<b>0.00</b>	<b>0.02</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>
FHC 1R $\mu$ -like	<b>0.62</b>	<b>0.12</b>	<b>0.16</b>	<b>0.02</b>	<b>0.03</b>	<b>0.00</b>	<b>0.03</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>
S.G. $\pi^0$ -like	<b>0.05</b>	<b>0.01</b>	<b>0.02</b>	<b>0.00</b>	<b>0.01</b>	<b>0.68</b>	<b>0.06</b>	<b>0.07</b>	<b>0.06</b>	<b>0.04</b>
S.G. $\mu$ -like 2de	<b>0.04</b>	<b>0.01</b>	<b>0.80</b>	<b>0.10</b>	<b>0.04</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>
S.G. $\mu$ -like 1de	<b>0.72</b>	<b>0.11</b>	<b>0.12</b>	<b>0.01</b>	<b>0.02</b>	<b>0.00</b>	<b>0.01</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>
S.G. $\mu$ -like 0de	<b>0.68</b>	<b>0.11</b>	<b>0.10</b>	<b>0.01</b>	<b>0.02</b>	<b>0.01</b>	<b>0.05</b>	<b>0.01</b>	<b>0.00</b>	<b>0.02</b>
S.G. e-like 1de	<b>0.05</b>	<b>0.01</b>	<b>0.75</b>	<b>0.10</b>	<b>0.05</b>	<b>0.00</b>	<b>0.01</b>	<b>0.02</b>	<b>0.00</b>	<b>0.01</b>
S.G. e-like 0de	<b>0.73</b>	<b>0.11</b>	<b>0.10</b>	<b>0.01</b>	<b>0.02</b>	<b>0.02</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>	<b>0.00</b>

**Figure 6.8.:** The interaction mode contribution of each sample given as a fraction of the total event rate in that sample. Asimov A oscillation parameters are assumed (given in Table 2.2). The Charged Current (CC) modes are broken into quasi-elastic (QE), 2p2h, resonant charged pion production ( $1\pi^\pm$ ), multi-pion production ( $M\pi$ ), and other interaction categories. Neutral Current (NC) interaction modes are given in interaction mode categories:  $\pi^0$  production, resonant charged pion production, multi-pion production, and others.

1989 which the Monte Carlo was generated. Secondly, the set of dial values which are taken  
 1990 from external data measurements and used as inputs. These are the “pre-fit” dial  
 1991 values. The reason these two sets of dial values are different is that the external data  
 1992 measurements are continually updated but it is very computationally intensive to  
 1993 regenerate a Monte Carlo prediction after each update. The final tune is the “post-fit”,  
 1994 “post-ND fit” or “post-BANFF” dial values. These are the values taken from a fit to the  
 1995 beam near detector data. This fit is performed by two independent fitting frameworks,  
 1996 MaCh3 and BANFF, which ensures reliable measurements. The output of each fitter is  
 1997 converted into a covariance matrix to describe the error and correlations between  
 1998 all the flux and cross-section parameters. This is then propagated to the far-detector

1999 oscillation analysis group for use in the P-Theta fitting framework. As MaCh3 can  
2000 perform a near detector fit, it is included within the simultaneous fit of the far-detector  
2001 beam and atmospheric oscillation analysis. This is because this technique does not  
2002 require any assumption of Gaussian posterior distributions which is required in the  
2003 covariance matrix methodology.

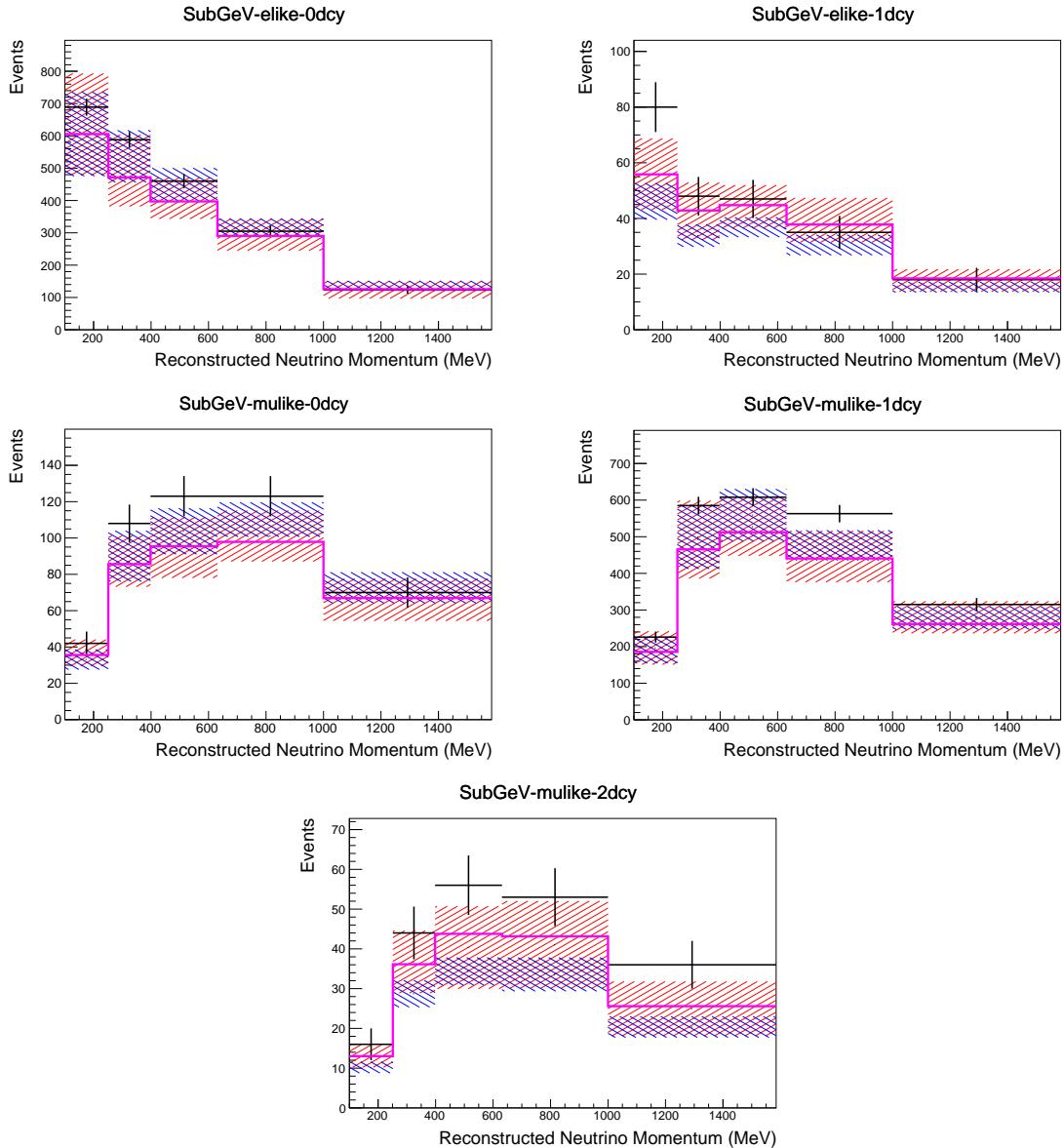
2004 On top of the combination of the SK and T2K interaction models, several other  
2005 parameters have been specifically developed for the joint oscillation analysis. The  
2006 majority of the atmospheric samples'  $\delta_{CP}$  sensitivity comes from the normalisation  
2007 of subGeV electron-like events. These are modeled using a spectral function model  
2008 to approximate the nuclear ground state. However, the near detector is not able to  
2009 constrain the model. Therefore, an additional systematic is introduced which models  
2010 an alternative Continous Random Phase Approximation (CRPA) nuclear ground  
2011 state. The reasoning is documented in [205]. As the near detector can not sufficiently  
2012 constrain the model, this dial approximates the event weights if a CRPA model had  
2013 been assumed rather than a spectral function. This dial only effects  $\nu_e$  and  $\bar{\nu}_e$  and is  
2014 applied as a shape parameter.

2015 Further additions to the model have been introduced due to the inclusion of the  
2016 subGeV  $\pi^0$  atmospheric sample. This particularly targets charged current and neutral  
2017 current  $\pi^0$  producing interactions to help constrain the systematic uncertainties. There  
2018 is no analogous sample in the T2K beam-only analysis so no significant effort has  
2019 been placed into building a sufficient uncertainty model. Therefore, an uncertainty  
2020 that affects neutral current resonant  $\pi^0$  production is incorporated into this analysis.  
2021 Comparisons of NEUT's NC resonant pion production predictions have been made to  
2022 MiniBooNE [206] data and a consistent 16% to 21% underprediction is observed [205].  
2023 Consequently, a conservative 30% normalisation parameter is invoked.

2024 Down-going events are mostly insensitive to oscillation parameters and can act  
2025 similar to the near detector within an accelerator experiment (Details will be dis-  
2026 cussed in chapter 7). This region of phase space can act as a sideband and allows the  
2027 cross-section model and near detector constraint to be studied. The distribution of  
2028 events in this region is calculated using the technique outlined in subsection 4.3.4. For  
2029 CCQE-targeting samples, the application of the near detector constraint is well within  
2030 the statistical fluctuation of the down-going data such that no significant tension is  
2031 observed between the data and the Monte Carlo prediction after the near detector  
2032 constraint is applied. This is not the case for samples with target CCRES interac-  
2033 tions. The electron-like data is consistent with the constrained prediction at high  
2034 reconstructed momenta but diverges at lower momentum, whereas the muon-like  
2035 sample is under-predicted throughout the range of momenta. To combat this disagree-  
2036 ment, an additional cross-section systematic dial, specifically designed to inflate the  
2037 low pion momentum systematics was developed in [205]. This is a shape parameter  
2038 implemented through a splined response.

#### 2039 6.4.4. Near Detector

2040 The systematics applied due to uncertainties arising from the response of the near  
2041 detector is documented in [133]. The response is described by 574 normalisation param-  
2042 eters binned in the selected sample as well as momentum and angle,  $P_\mu$  and  $\cos(\theta_\mu)$ ,  
2043 of the final-state muon. These are applied via a covariance matrix with each parameter  
2044 being assigned a Gaussian prior from that covariance matrix. These normalisation  
2045 parameters are built from underlying systematics, e.g. pion secondary interaction  
2046 systematics, which are randomly thrown and the variation in each  $P_\mu \times \cos(\theta_\mu)$  bin is  
2047 determined. Two thousand throws are evaluated and a covariance matrix response is  
2048 created. This allows significant correlations between FGD1 and FGD2 samples, as well



**Figure 6.9.:** Down-going atmospheric subGeV single-ring samples comparing the mean and error of the pre-fit and post-fit Monte Carlo predictions in red and blue, respectively. The magenta histogram illustrates the Monte Carlo prediction using the generated dial values. The black points illustrate the down-going data with statistical errors given. The mean and errors of the Monte Carlo predictions are calculated by the techniques documented in subsection 4.3.4. The pre-fit spectrum is calculated by throwing the cross-section and atmospheric flux dial values from the pre-fit covariance matrix. The post-fit spectrum is calculated by sampling the cross-section dial values from an ND fit MCMC chain, whilst still throwing the atmospheric flux dials from the pre-fit covariance.

as adjacent  $P_\mu \times \cos(\theta_\mu)$  bins. Statistical uncertainties are accounted for by including fluctuations of each event's weight from a Poisson distribution.

2051 Similar to the cross-section systematics, MaCh3 and BANFF are used to constrain  
2052 the uncertainty of these systematics through independent validations. Each fitter  
2053 generates a post-fit covariance matrix which is compared and passed to the far-detector  
2054 oscillation analysis working group. As the analysis presented within this thesis uses  
2055 the MaCh3 framework, a joint oscillation analysis fit of all three sets of samples and  
2056 their respective systematics is performed.

2057 **6.4.5. Far Detector**

2058 Two configurations of the far detector systematic model implementation have been  
2059 considered. Firstly, the far detector systematic uncertainties for beam and atmospheric  
2060 samples are taken from their respective analysis inputs, denoted “official inputs” anal-  
2061 ysis. Consequently, no correlations are assumed between the beam and atmospheric  
2062 samples. The generation of the beam- and atmospheric-specific inputs are documented  
2063 in subsubsection 6.4.5.1 and subsubsection 6.4.5.2. Secondly, a correlated detector  
2064 model has been developed. Here, the distribution of parameters used for applying  
2065 event cuts (e.g. electron-muon PID separation) is modified within the fit. It follows  
2066 a similar methodology to the beam far detector systematics implementation but per-  
2067 forms a joint fit of the beam and atmospheric data. This alternative implementation is  
2068 detailed in subsubsection 6.4.5.3.

2069 **6.4.5.1. Beam Samples**

2070 There are 45 systematics which describe the response of the far detector to beam  
2071 events [185], split into 44 normalisation parameters and one energy scale systematic.  
2072 The energy scale systematic is applied as a multiplicative scaling of the reconstructed  
2073 neutrino energy. It is described by a Gaussian, centered at one with equal to the

2074 difference in Monte Carlo to data comparisons performed in [187]. The normalisation  
 2075 parameters are assigned a Gaussian error centralised at one with width taken from a  
 2076 covariance matrix. A detailed breakdown of the generation of the covariance matrix  
 2077 is found in [199]. To build the covariance matrix, a fit is performed on atmospheric  
 2078 data which has been selected using beam sample selection cuts. These cuts use the  
 2079 variables,  $L^i$ , where the index  $i$  is detailed in Table 6.6. Each  $L^i$  is a smear,  $\alpha$ , and shift,  
 2080  $\beta$  parameter such that,

$$L_j^i \rightarrow \bar{L}_j^i = \alpha_j^i L + \beta_j^i \quad (6.4)$$

2081 Where  $L_j^i$  ( $\bar{L}_j^i$ ) correspond to nominal(varied) PID cut parameters given in Table 6.6.  
 2082 The shift and smear parameters are binned by final-state topology,  $j$ , where the binning  
 2083 is given in Table 6.7. The final-state topology binning is because the detector will  
 2084 respond differently to events that have one or multiple rings. For example, the detector  
 2085 will be able to distinguish single-ring events better than two overlapping ring events,  
 2086 resulting in smaller systematic uncertainty for one-ring events compared to two-ring  
 2087 events. This approach is used to allow the cut parameter distributions to be modified  
 2088 within the fit, allowing for better data to Monte Carlo agreement. Only the shape  
 2089 of each of the cut variables is used within this fit, such that physics effects are not  
 2090 considered.

Cut Variable	Parameter Name
0	<code>fitQun e/mu PID</code>
1	<code>fitQun e/pi0 PID</code>
2	<code>fitQun mu/pi PID</code>
3	<code>fitQun Ring-Counting Parameter</code>

**Table 6.6.:** List of cut variables that are included within the shift/smear fit documented in [199].

Category	Description
$1e$	Only one electron above Cherenkov threshold in the final state
$1\mu$	Only one muon above Cherenkov threshold in the final state
$1e+other$	One electron and one or more other charged particles above Cherenkov threshold in the final state
$1\mu+other$	One muon and one or more other charged particles above Cherenkov threshold in the final state
$1\pi^0$	Only one $\pi^0$ in the final state
$1\pi^\pm$ or $1p$	Only one hadron (typically charged pion or proton) in the final state
Other	Any other final state

**Table 6.7.:** Reconstructed event topology categories on which the SK detector systematics [199] are based.

2091 Beyond the uncertainty on the PID cut criteria, the mis-modeling of  $\pi^0$  events  
 2092 is also considered. If one of the two rings from a  $\pi^0$  event is missed, this will be  
 2093 reconstructed as a  $CC\nu_e$ -like event. This is one of the largest systematics hindering the  
 2094 electron neutrino appearance analyses. Consequently, additional systematics has been  
 2095 introduced to constrain the mis-modeling of  $\pi^0$  events in SK, binned by reconstructed  
 2096 neutrino energy. To evaluate this systematic uncertainty, a set of “hybrid- $\pi^0$ ” samples  
 2097 is constructed. These events are built by overlaying one electron-like ring from the  
 2098 SK atmospheric neutrino samples or decay electron ring from a stopping cosmic ray  
 2099 muon with one simulated photon ring. Both rings are chosen so that momenta and  
 2100 opening angle follow the decay kinematics of NC  $\pi^0$  events from the T2K-MC. Hybrid-  
 2101  $\pi^0$  Monte Carlo samples with both rings from the SK Monte Carlo are produced  
 2102 to compare with the hybrid- $\pi^0$  data samples and the difference in the fraction of  
 2103 events that pass the  $\nu_e$  selection criteria is used to assign the systematic errors. In  
 2104 order to investigate any data to Monte Carlo differences that may originate from  
 2105 either the higher energy ring or lower energy ring, two samples are built; a sample  
 2106 in which the electron constitutes the higher energy ring from the  $\pi^0$  decay called the  
 2107 primary sample and another one in which it constitutes the lower energy ring called

2108 the secondary sample. The standard T2K  $\nu_e$  fiTQun event selection criteria are used to  
2109 select events.

2110 Final contributions to the covariance matrix are determined by supplementary  
2111 uncertainties attained by comparing stopping muon data to Monte Carlo prediction,  
2112 as first introduced in section 5.2. The efficiency of tagging decay electrons is estimated  
2113 by the stopping muon data to Monte Carlo differences by comparing the number  
2114 of one decay electron events to the number of events with one or fewer decay elec-  
2115 trons. Similarly, the rate at which fake decay electrons are reconstructed by fiTQun  
2116 is estimated by comparing the number of two decay electron events to the number  
2117 of events with one or two reconstructed decay electrons. The two sources of sys-  
2118 tematics are added in quadrature weighted by the number of events with one true  
2119 decay electron yielding a 0.2% systematic uncertainty. A fiducial volume systematic of  
2120  $\pm 2.5\text{cm}$  which corresponds to a 0.5% shift in the normalisation of events. Additional  
2121 normalisation uncertainties based on neutrino flavour and interaction mode are also  
2122 defined in [185, 207, 208].

2123 Two additional sources of uncertainty are included: secondary and photo-nuclear  
2124 interactions. These are estimated by varying the underlying parameters are building a  
2125 distribution of sample event rates. These contributions are then added in quadrature  
2126 to the above covariance matrix.

#### 2127 6.4.5.2. Atmospheric Samples

2128 The systematic parameters which control the detector systematics for atmospheric  
2129 samples, documented in [89], are split into two sub-groups. Those which are related  
2130 to particle identification and ring counting systematics and those which are related to  
2131 calibration, separation, and reduction uncertainties.

2132        The particle identification systematics consist of five parameters. The ring separation systematic enforces an anti-correlated response between the single-ring and  
2133        multi-ring samples. This is implemented as a fractional increase/decrease in the  
2134        overall normalisation of each sample, depending on the distance to the nearest wall  
2135        from an event's vertex. The coefficients of the normalisation are estimated prior to the  
2136        fit and depend on the particular atmospheric sample. The single-ring and multi-ring  
2137        PID systematics encode the detector's ability to separate electron-like and muon-like  
2138        events and are implemented in a similar way to the ring separation systematic.  
2139

2140        The multi-ring electron-like separation likelihood, discussed in section 6.1, encodes  
2141        the ability of the detector to separate neutrino from anti-neutrino events. As an impor-  
2142        tant systematic in the mass hierarchy determination, systematic uncertainties control  
2143        the relative normalisations of the  $\nu_e$  and  $\bar{\nu}_e$  enriched samples. Two normalisation  
2144        parameters are implemented which vary the event rate of each multi-ring sample,  
2145        whilst ensuring the total event rate is conserved.

2146        There are 22 systematics related to calibration measurements, including effects  
2147        from backgrounds, reduction, and showering effects. They are documented in [89] and  
2148        briefly summarised in Table 6.8. They are applied via normalisation parameters, with  
2149        the separation systematics requiring the conservation of event rate across all samples.

2150        **6.4.5.3. Correlated Detector Model**

2151        A complete uncertainty model of the SK detector would be able to determine the  
2152        systematic shift on the sample spectra for a variation of the underlying parameters,  
2153        e.g. PMT angular acceptance. However, this is computationally intensive, requiring  
2154        Monte Carlo predictions to be made for each plausible variation. Consequently, an  
2155        effective parameter model has been utilised for a correlated detector model following  
2156        from the T2K-only model implementation documented in subsubsection 6.4.5.1. The

**Table 6.8.:** Sources of systematic errors specified within the grouped into the “calibration” systematics model.

Index	Description
0	Partially contained reduction
1	Fully contained reduction
2	Separation of fully contained and partially contained events
3	Separation of stopping and through-going partially contained events in top of detector
4	Separation of stopping and through-going partially contained events in barrel of detector
5	Separation of stopping and through-going partially contained events in bottom of detector
6	Background due to cosmic rays
7	Background due to flasher events
8	Vertex systematic moving events into and out of fiducial volume
9	Upward going muon event reduction
10	Separation of stopping and through-going in upward going muon events
11	Energy systematic in upward going muon events
12	Reconstruction of the path length of upward going muon events
13	Separation of showering and non-showering upward going muon events
14	Background of stopping upward going muon events
15	Background of non-showering through-going upward going muon events
16	Background of showering through-going upward going muon events
17	Efficiency of tagging two rings from $\pi^0$ decay
18	Efficiency of decay electron tagging
19	Background from downgoing cosmic muons
20	Asymmetry of energy deposition in tank
21	Energy scale deposition

<sup>2157</sup> implementation performs a simultaneous fit of detector and oscillation parameters,  
<sup>2158</sup> for the detector parameters given in Table 6.6.

<sup>2159</sup> The correlated detector model utilises the same smear and shift parameters docu-  
<sup>2160</sup> mented in subsubsection 6.4.5.1, split by final state topology. Beyond this, the shift  
<sup>2161</sup> and smear parameters are split by visible energy deposited within the detector, with

2162 binning specified in Table 6.9. This is because atmospheric events are categorised by  
2163 subGeV and multiGeV events based on visible energy, so this splitting is required  
2164 when correlating the systematic model for beam and atmospheric events. Alongside  
2165 the technical requirement, higher energy events will be better reconstructed due to  
2166 fractionally less noise within the detector. This implementation correlates the detector  
2167 systematics between the far-detector beam and subGeV atmospheric samples due  
2168 to their similar energies and interaction types. As a result of the inclusion of visible  
2169 energy binning, Equation 6.4 becomes

$$L_{jk}^i \rightarrow \bar{L}_{jk}^i = \alpha_{jk}^i L + \beta_{jk}^i, \quad (6.5)$$

2170 where  $k$  is the visible energy bin. As there are no equivalent beam samples, the  
2171 multiGeV, multiring, PC, and Up- $\mu$  samples will be subject to the ATMPD particle  
2172 identification systematics implementation as described in subsubsection 6.4.5.2 rather  
2173 than using this correlated detector model. The calibration systematics also described  
2174 in the aforementioned chapter still apply to all atmospheric samples.

Index	Range (MeV)
0	$30 \geq x > 300$
1	$300 \geq x > 700$
2	$700 \geq x > 1330$
3	$1330 \geq x$

**Table 6.9.:** Visible energy binning for which the correlated SK detector systematics are based

2175 The implementation of this systematic model takes the events reconstructed values  
2176 of the cut parameters, modifies them by the particular shift and smear parameter for  
2177 that event, and then re-applies event selection. This invokes event migration, which is

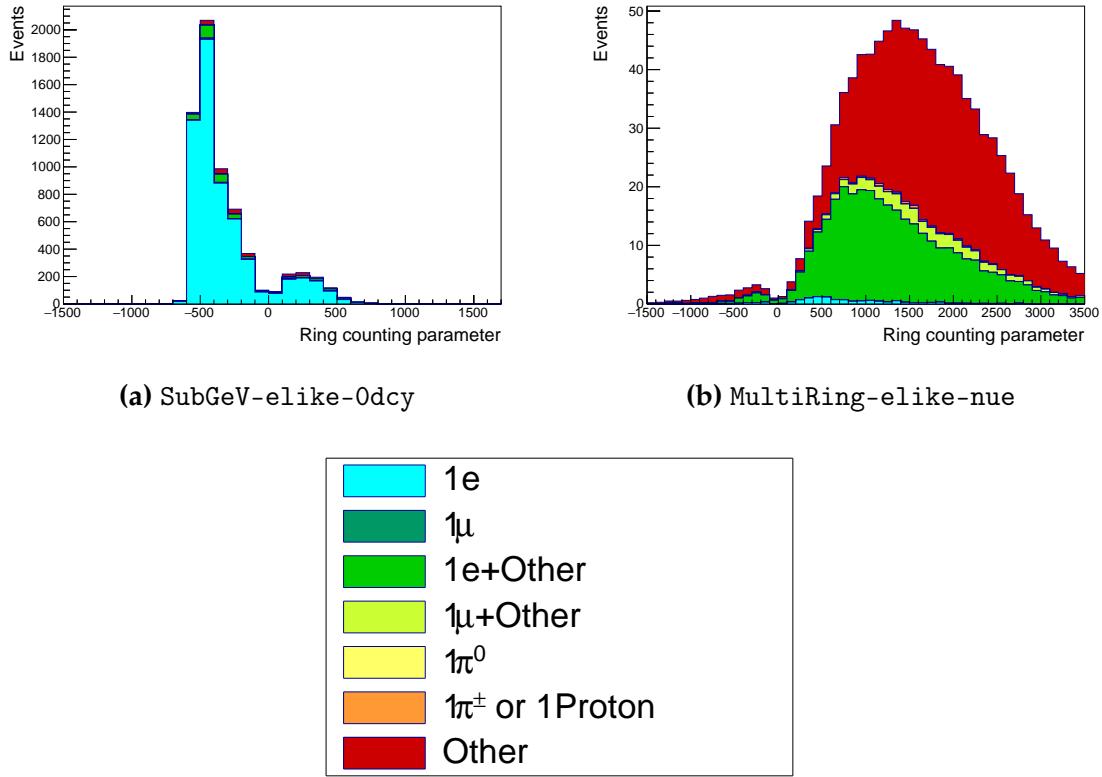
2178 a new feature incorporated into the MaCh3 framework which is only achievable due  
2179 to the event-by-event reweighting scheme.

2180 Particular care has to be taken when varying the ring counting parameter. This  
2181 is because the number of rings is a finite value (one-ring, two-rings, etc.) which can  
2182 not be continuously varied through this shift and smear technique. Consequently a  
2183 ring counting parameter,  $RC_i$ , is calculated for the  $i^{th}$  event, following the definition  
2184 in [184]. The likelihood from all considered one-ring ( $L_{1R}$ ) and two-ring ( $L_{2R}$ ) fits  
2185 are compared to determine the preferred hypothesis. The difference is computed as  
2186  $\Delta_{LLH} = \log(L_{1R}) - \log(L_{2R})$ . The ring counting parameter is then defined as,

$$RC_i = \text{sgn}(\Delta_{LLH}) \times \sqrt{|\Delta_{LLH}|}, \quad (6.6)$$

2187 where  $\text{sgn}(x) = x/|x|$ . This ring counting parameter corresponds to an intermedi-  
2188 ate likelihood value used within the `fitQun` algorithm to decide the number of rings  
2189 associated with a particular event. However, fake-ring merging algorithms are applied  
2190 after this likelihood value is used. Consequently, this ring counting parameter does  
2191 not always exactly correspond to the number of reconstructed rings. This can be seen  
2192 in Figure 6.10.

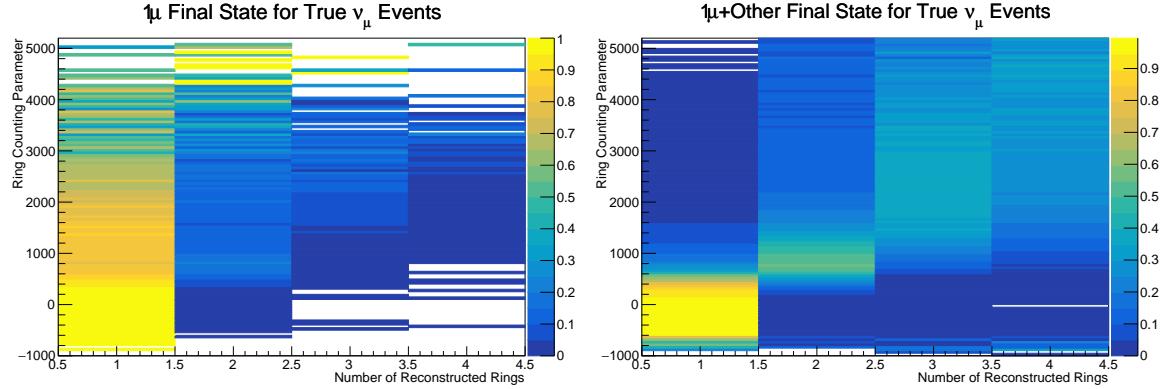
2193 As the `fitQun` algorithm does not provide a likelihood value after the fake-ring  
2194 algorithms have been applied, the ring counting parameter distribution is correlated to  
2195 the final number of reconstructed rings through “maps”. These are two-dimensional  
2196 distributions of the ring counting parameter and the final number of reconstructed  
2197 rings. An example is illustrated in Figure 6.11. In principle, the `fitQun` reconstruction  
2198 algorithm should be re-run after the variation in the ring counting parameter. However,



**Figure 6.10.:** The ring counting parameter as defined in Equation 6.6 for the SubGeV-else-0dcy and MultiRing-else-nue samples.

2199 this is not computationally viable. Therefore the “maps” are used as a reweighting  
 2200 template.

2201 The maps are split by final state topology and true neutrino flavour and all `fitQun`  
 2202 -reconstructed Monte Carlo events are used to fill them. The maps are row-normalised  
 2203 to represent the probability of X number of rings for a given  $RC_i$  value. Prior to the  
 2204 fit, an event’s nominal weight is calculated as  $W^i(N_{Rings}^i, L_{jk}^i)$ , where  $N_{Rings}^i$  is the  
 2205 reconstructed number of rings for the  $i^{th}$  event and  $W^i(x, y)$  is the bin content in map  
 2206 associated with the  $i^{th}$  event, where  $x$  number of rings and  $y$  is ring counting parameter.  
 2207 Then during the fit, the value of  $R = W^i(N_{Rings}^i, \bar{L}_{jk}^i) / W^i(N_{Rings}^i, L_{jk}^i)$  is calculated as  
 2208 the event weight for the  $i^{th}$  event. This is the only cut variable that uses a reweighting  
 2209 scheme rather than event migration.

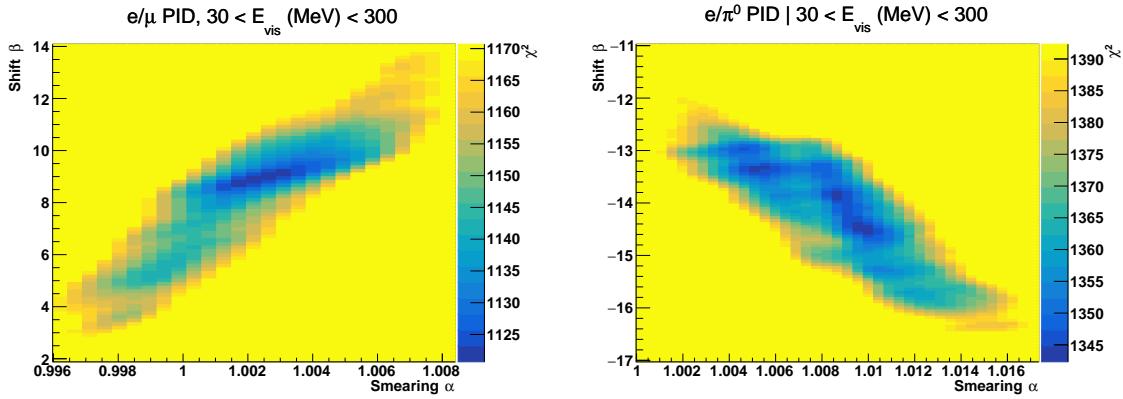


**Figure 6.11.:** The ring counting parameter, defined in Equation 6.6, as a function of the number of reconstructed rings as found by the `fitQun` reconstruction algorithm. Left: true  $\nu_\mu$  events with only one muon above the Cherenkov threshold in the final state. Right: true  $\nu_\mu$  events with one muon and at least one other charged particle above the Cherenkov threshold in the final state.

2210     The  $\pi^0$  systematics introduced in subsection 6.4.4 are applied via a covariance  
 2211     matrix. This is not possible in the alternative model as no covariance matrix is used.  
 2212     Thus, the implementation of the  $\pi^0$  systematics has been modified. The inputs from  
 2213     the hybrid  $\pi^0$  sample is included via the use of “ $\chi^2$  maps”, which are two-dimensional  
 2214     histograms in  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters over some range. Illustrative examples of the  $\chi^2$   
 2215     maps are given in Figure 6.12. Due to their nature, the shift and smear parameters are  
 2216     typically very correlated. A map is produced for each cut parameter given in Table 6.6  
 2217     and for each visible energy bin given in Table 6.9.

2218     The maps are filled through the  $\chi^2$  comparison of the hybrid  $\pi^0$  Monte Carlo and  
 2219     data in the particle identification parameters documented in Table 6.6. The Monte  
 2220     Carlo distribution is modified by the  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  scaling, whilst cross-section and flux  
 2221     nuisance parameters are thrown from their prior uncertainties. The  $\chi^2$  between the  
 2222     scaled Monte Carlo and data is calculated and the relevant point in the  $\chi^2$  map is filled.

2223     The implementation within this alternative detector model is to add the bin contents  
 2224     of the maps, for the relevant values of the  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters, to the likelihood  
 2225     penalty. Only  $1\pi^0$  final state topology shift and smear parameters use this prior  
 2226     uncertainty.



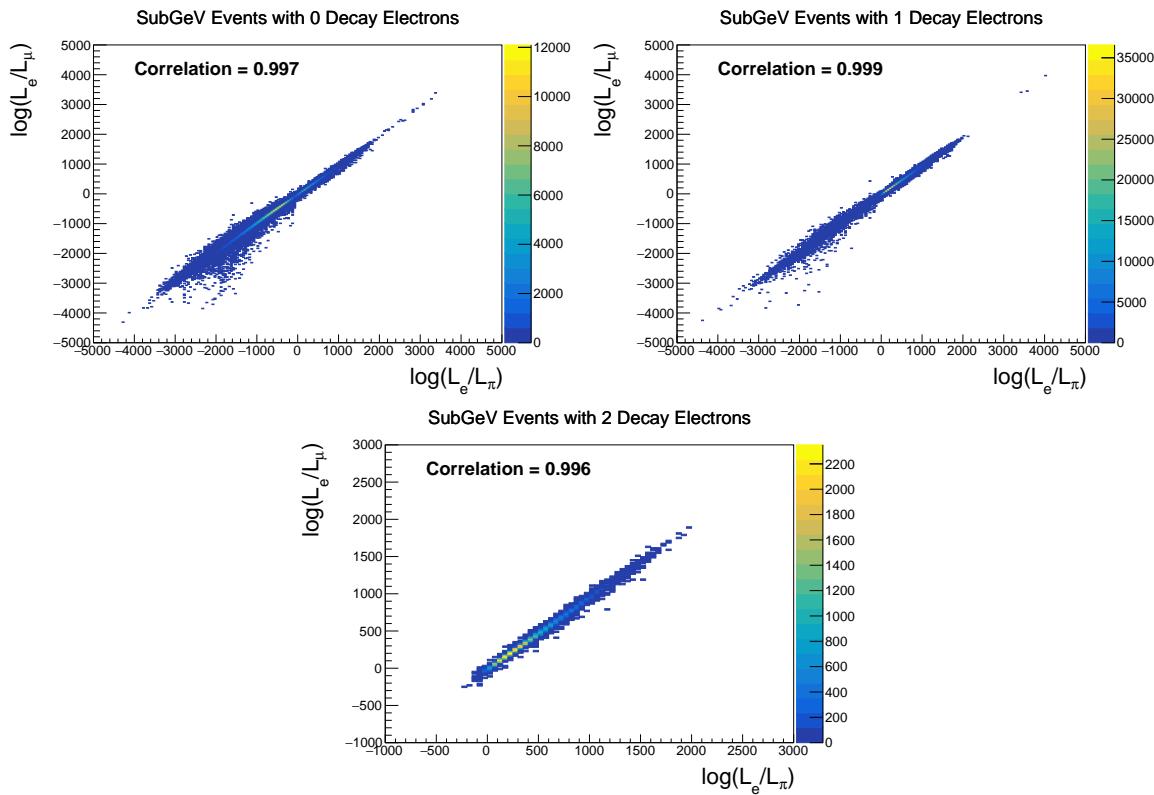
**Figure 6.12.:** The  $\chi^2$  between the hybrid- $\pi^0$  Monte Carlo and data samples, as a function of smear ( $\alpha$ ) and shift ( $\beta$ ) parameters, for events which have  $1\pi^0$  final state topology. Left: Electron-muon separation PID parameter for events with  $30 \geq E_{vis}(MeV) < 300$ . Right: Electron- $\pi^0$  separation PID parameter for events with  $30 \geq E_{vis}(MeV) < 300$ .

2227 Similarly, the implementation of the supplementary systematics documented in  
 2228 subsubsection 6.4.5.1 needs to be modified. A new framework [209] was built in  
 2229 tandem with the T2K-SK working group [185] so the additional parameters can be  
 2230 incorporated into the MaCh3 framework. These are applied as normalisation parame-  
 2231 ters, depending on the particular interaction mode, number of tagged decay electrons,  
 2232 and whether the primary particle generated Cherenkov light. They are assigned  
 2233 Gaussian uncertainties with widths described by a covariance matrix. Furthermore,  
 2234 the secondary interaction and photo-nuclear effects need to be accounted for in this  
 2235 detector model using a different implementation than that in subsubsection 6.4.5.1.  
 2236 This was done by including a shape parameter for each of the secondary interactions  
 2237 and the photo-nuclear systematic parameters.

2238 There are a total of 224  $\alpha_{jk}^i$  and  $\beta_{jk}^i$  parameters, of which 32 have prior constraints  
 2239 from the hybrid  $\pi^0$  samples.

2240 One final complexity of this correlated detector model is that the two sets of sam-  
 2241 ples, beam and subGeV atmospheric, use slightly different parameters to distinguish  
 2242 electron and muon-like events. The T2K samples use the value of  $\log(L_e/L_\mu)$  whereas  
 2243 the atmospheric samples use the value of  $\log(L_e/L_\pi)$ , where  $L_X$  is the likelihood for

2244 hypothesis X. This is because the T2K fits use single-ring `fitTQu` fitting techniques,  
 2245 whereas multi-ring fits are applied to the atmospheric samples where only the electron  
 2246 and pion hypothesis are considered. The correlation between the two likelihood ratios  
 2247 is illustrated in Figure 6.13. As discussed in section 5.2, the pion hypothesis is a very  
 2248 good approximation of the muon hypothesis due to their similar mass. Consequently,  
 2249 using the same shift and smear parameters correlated between the beam and subGeV  
 2250 atmospheric samples is deemed a good approximation.



**Figure 6.13.:** The distribution of  $\log(L_e/L_\mu)$  compared to  $\log(L_e/L_\pi)$  for subGeV events with zero (top left), one (top right) or two (bottom) decay electrons. The correlation in the distribution is calculated as 0.997, 0.999 and 0.996, respectively.

2251

# Chapter 7

2252

## Oscillation Probability Calculation

2253 It is important to understand how and where the sensitivity to the oscillation pa-  
2254 rameters comes from for both atmospheric and beam samples. An overview of how  
2255 these samples observe changes in  $\delta_{CP}$ ,  $\Delta m_{23}^2$ , and  $\sin^2(\theta_{23})$  is given in section 7.1. It  
2256 also explains the additional complexities involved when performing an atmospheric  
2257 neutrino analysis as compared to a beam-only analysis.

2258 Without additional techniques, atmospheric sub-GeV upward-going neutrinos  
2259 ( $E_\nu < 1.33\text{GeV}, \cos(\theta_Z) < 0.$ ) can artificially inflate the sensitivity to  $\delta_{CP}$  due to the  
2260 quickly varying oscillation probability in this region. Therefore, a “sub-sampling”  
2261 approach has been developed to reduce these biases ensuring accurate and reliable  
2262 sensitivity measurements. This technique ensures that small-scale unresolvable fea-  
2263 tures of the oscillation probability have been averaged over whilst the large-scale  
2264 features in the oscillation probability are unaffected. The documentation and valida-  
2265 tion of this technique are found in section 7.2. The oscillation probability calculation is  
2266 computationally intensive due to the large number of matrix multiplications needed.  
2267 Consequently, the CUDAProb3 implementation choice made within the fitting frame-  
2268 work, as detailed in section 7.3, ensures that the analysis can be done in a timely  
2269 manner.

2270 Whilst the beam neutrinos are assumed to propagate through a constant density  
2271 slab of material, the density variations through the Earth result in more complex  
2272 oscillation patterns. Furthermore, the uncertainty in the electron density can modify  
2273 the oscillation probability for the denser core layers of the Earth. The model of the

2274 Earth used within this analysis is detailed in section 7.4. This includes information  
2275 about the official SK-only methodology as well as improvements that can be made  
2276 to remove some of the approximations made in that analysis. Another complexity of  
2277 atmospheric neutrinos oscillation studies is that the height of production in the atmo-  
2278 sphere is not known on an event-by-event basis. An analytical averaging technique  
2279 that approximates the uncertainty of the oscillation probability has been followed,  
2280 with the author of this thesis being responsible for the implementation and validation.  
2281 This implementation of an external technique is illustrated in section 7.5.

## 2282 7.1. Overview

2283 DB: Should this be moved into an earlier chapter? The selections chapter references  
2284 the matter resonance which has not yet been explained at that point

2285 The analysis presented within this thesis focuses on the determination of oscillation  
2286 parameters from atmospheric and beam neutrinos. Whilst subject to the same oscil-  
2287 lation formalism, the way in which the two samples have sensitivity to the different  
2288 oscillation parameters differs quite significantly.

2289 Atmospheric neutrinos have a varying baseline, or “path length”,  $L$ , such that  
2290 the distance each neutrino travels before interacting is dependent upon the zenith  
2291 angle,  $\theta_Z$ . As primary cosmic rays can interact anywhere between the Earth’s surface  
2292 and  $\sim 50\text{km}$  above that, the height,  $h$ , in the atmosphere at which the neutrino was  
2293 generated also affects the path length,

$$L = \sqrt{(R_E + h)^2 - R_E^2 (1 - \cos^2(\theta_Z))} - R_E \cos(\theta_Z). \quad (7.1)$$

2294 Where  $R_E = 6,371\text{km}$  is the Earth's radius. Consequently, the oscillation probabil-  
2295 ity is dependent upon two parameters,  $\cos(\theta_Z)$  and  $E_\nu$ .

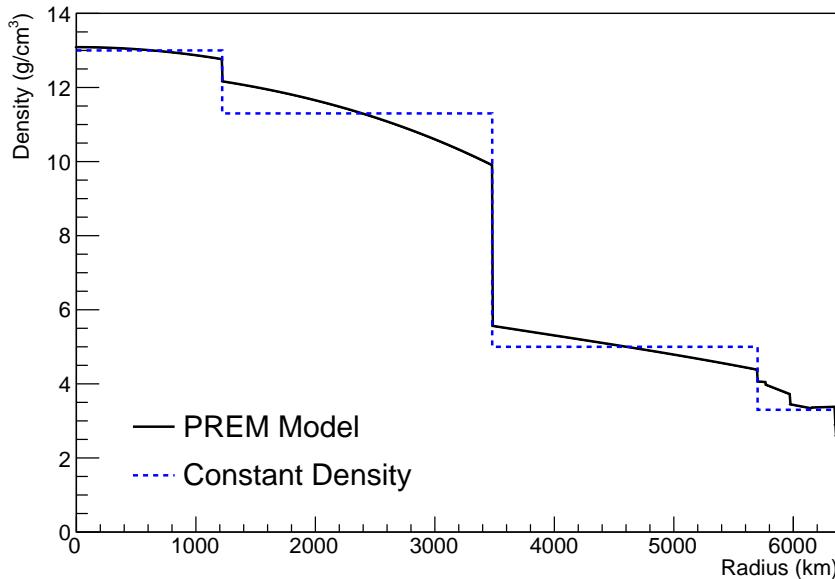
2296 The oscillation probability used within this analysis is based on [21]. The neutrino  
2297 wavefunction in the vacuum Hamiltonian evolves in each layer of constant matter  
2298 density via

$$i \frac{d\psi_j(t)}{dt} = \frac{m_j^2}{2E_\nu} \psi_j(t) - \sum_k \sqrt{2} G_F N_e U_{ej} U_{ke}^\dagger \psi_k(t), \quad (7.2)$$

2299 where  $m_j^2$  is the square of the  $j^{th}$  vacuum eigenstate mass,  $E_\nu$  is the neutrino  
2300 energy,  $G_F$  is Fermi's constant,  $N_e$  is the electron number density and  $U$  is the PMNS  
2301 matrix. The transformation  $N_e \rightarrow -N_e$  and  $\delta_{CP} \rightarrow -\delta_{CP}$  is applied for antineutrino  
2302 propagation. Thus, a model of the Earth's density is required for atmospheric neutrino  
2303 propagation. Following the official SK-only methodology [210], this analysis uses the  
2304 Preliminary Reference Earth Model (PREM) [211]. This model provides piecewise cubic  
2305 polynomials as a function of the Earth's radius which results in the density profile  
2306 illustrated in Figure 7.1. As discussed, the propagator requires layers of constant  
2307 density. The SK methodology approximates the PREM model by using four layers of  
2308 constant density [210]. The details of these layers are detailed in Table 7.1.

Layer	Outer Radius [km]	Density [ $\text{g}/\text{cm}^3$ ]	Chemical composition (Z/A)
Inner Core	1220	13	$0.468 \pm 0.029$
Outer Core	3480	11.3	$0.468 \pm 0.029$
Lower Mantle	5701	5.0	0.496
Transition Zone	6371	3.3	0.496

**Table 7.1.:** Description of the four layers of the Earth invoked within the constant density approximation of the PREM model [211].

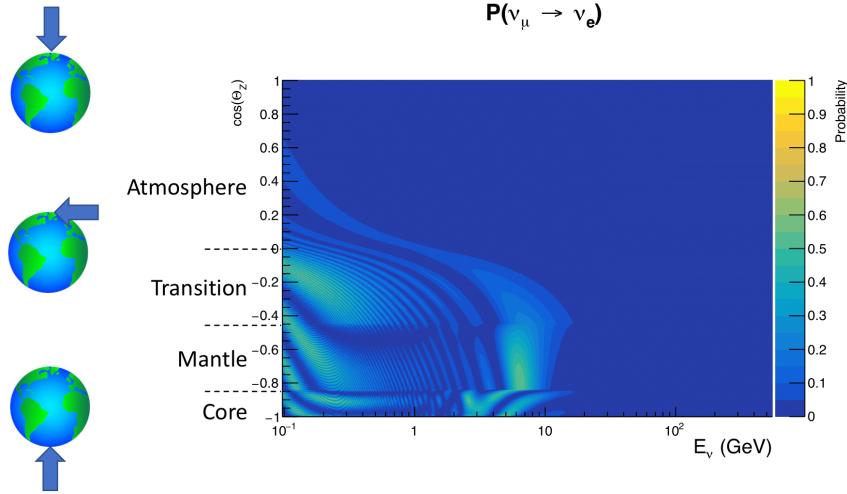


**Figure 7.1.:** The density of the Earth given as a function of the radius, as given by the PREM model (Black), and the constant density four-layer approximation (Blue), as used in the official SK-only analysis.

2309     The atmospheric neutrino oscillation probabilities can be presented as two dimen-  
 2310     sional “oscillograms” as illustrated in Figure 7.2. The distinct discontinuities, as a  
 2311     function of  $\cos(\theta_Z)$ , are due to the discrete change in density invoked within the PREM  
 2312     model.

2313     Atmospheric neutrinos do have sensitivity to  $\delta_{CP}$  through a normalisation term.  
 2314     Figure 7.3 illustrates the difference in oscillation probability between CP-conserving  
 2315     ( $\delta_{CP} = 0.$ ) and a CP-violating ( $\delta_{CP} = -1.601$ ) value taken from Asimov A oscillation  
 2316     parameter set (Table 2.2). The result is a complicated oscillation pattern in the appear-  
 2317     ance probability for sub-GeV upgoing neutrinos. The detector does not have sufficient  
 2318     resolution to resolve these individual patterns so the sensitivity to  $\delta_{CP}$  for atmospheric  
 2319     neutrinos comes via the overall normalisation of these events.

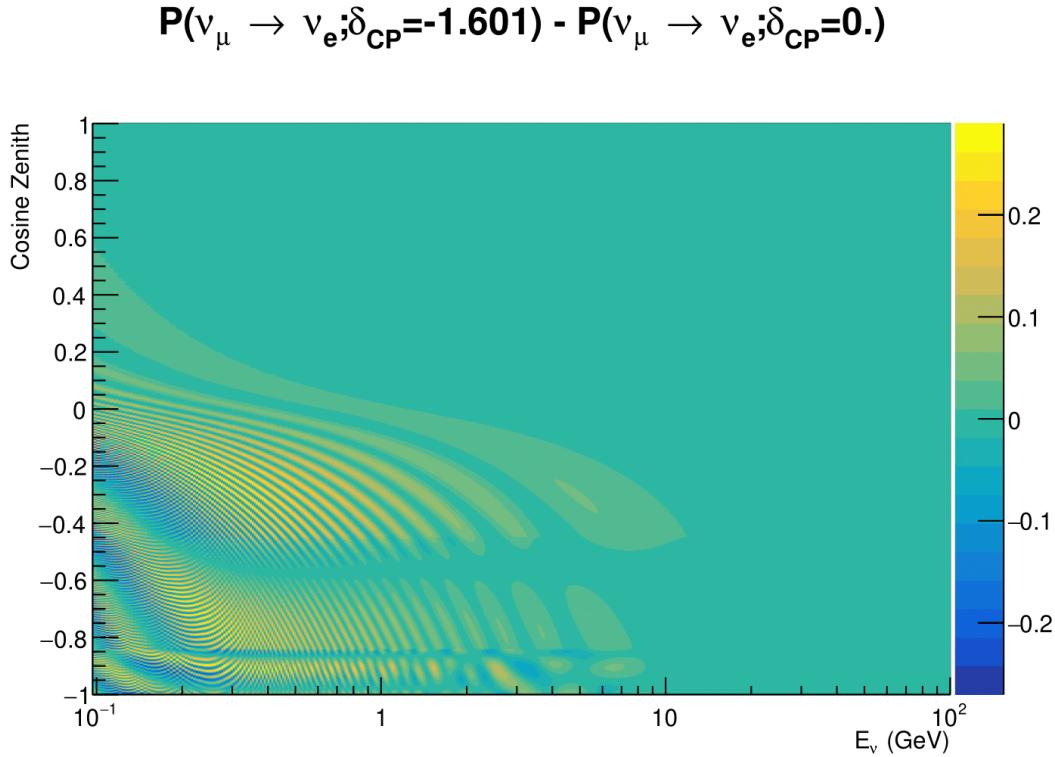
2320     The presence of matter means that the effect  $\delta_{CP}$  has on the oscillation probability  
 2321     is not equal between neutrinos and antineutrinos, which would be expected when  
 2322     propagating through a vacuum. This is further extenuated by the fact that SK can



**Figure 7.2.:** An “oscillogram” that depicts the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability as a function of neutrino energy and cosine of the zenith angle. The zenith angle is defined such that  $\cos(\theta_Z) = 1.0$  represents neutrinos that travel from directly above the detector. The four-layer constant density PREM model approximation is used and Asimov A oscillation parameters are assumed (Table 2.2).

not distinguish neutrinos and antineutrinos well and that the cross-section neutrino interaction is larger than that for antineutrinos. Finally, sample selections (discussed in section 6.1) targeting different neutrino interaction modes result in an imbalance in the percentage of neutrinos to anti-neutrinos. This is because negatively charged pions from antineutrino interactions are more likely to be captured by a nucleus compared to a positively charged pion. All of these effects lead to a difference in the number of neutrinos detected compared to antineutrinos. This changes how the  $\delta_{CP}$  normalisation term is observed, resulting in a very complex sensitivity to  $\delta_{CP}$ .

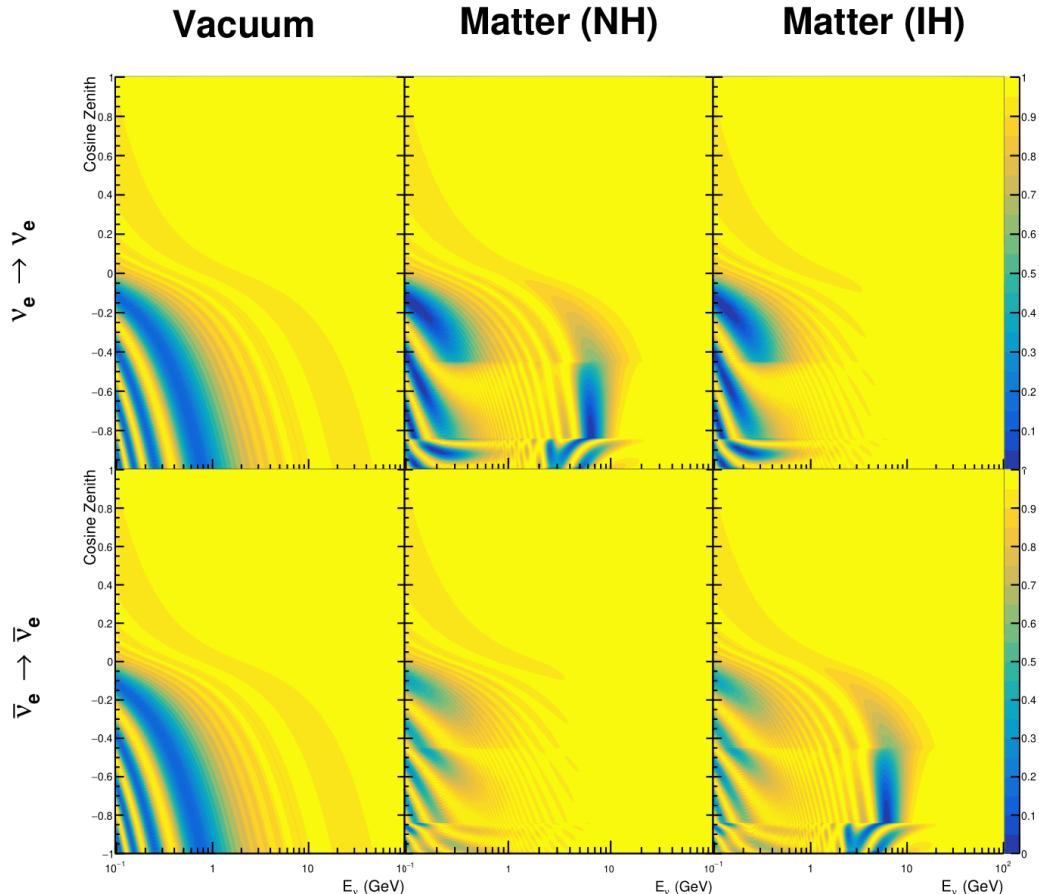
Atmospheric neutrinos are subject to matter effects as they travel through the dense matter in the Earth. The vacuum and matter oscillation probabilities for  $P(\nu_e \rightarrow \nu_e)$  and  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  are presented in Figure 7.4, where the PREM model has been assumed. The oscillation probability for both neutrinos and antineutrinos is affected in the presence of matter. However, the resonance effects around  $O(5)\text{GeV}$  only occur for neutrinos in normal mass hierarchy and antineutrinos in inverse mass hierarchy. The



**Figure 7.3.:** The effect of  $\delta_{CP}$  for atmospheric neutrinos given in terms of the neutrino energy and zenith angle. This oscillogram compares the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability for a CP conserving ( $\delta_{CP} = 0.0$ ) and a CP violating ( $\delta_{CP} = -1.601$ ) value taken from the Asimov A parameter set. The other oscillation parameters assume the Asimov A oscillation parameter set given in Table 2.2.

exact position and amplitude of the resonance depend on  $\sin^2(\theta_{23})$  meaning that the atmospheric neutrinos have sensitivity to  $\sin^2(\theta_{23})$ .

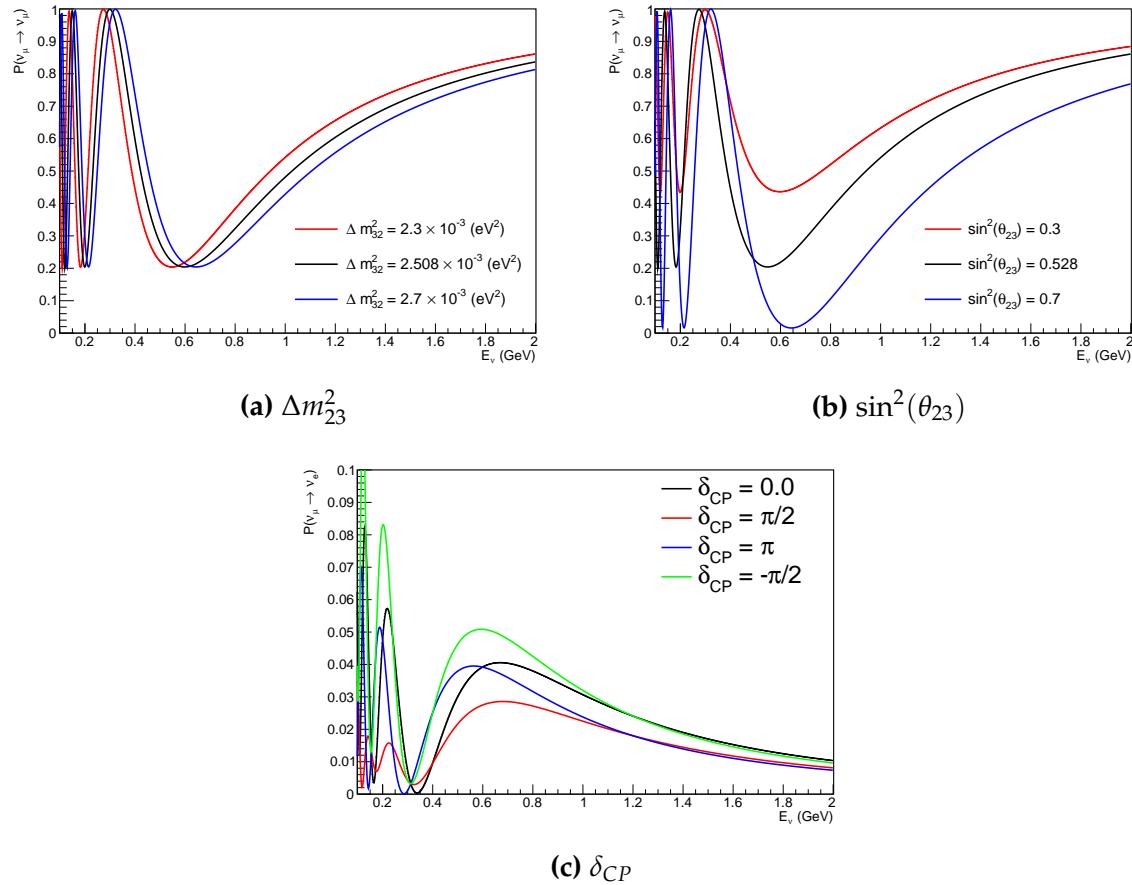
As the T2K beam flux is centered at the first oscillation maximum ( $E_\nu = 0.6\text{GeV}$ ), the sensitivity to  $\delta_{CP}$  is predominantly observed as a change in the event-rate of e-like samples in  $\nu/\bar{\nu}$  modes. Figure 7.5 illustrates the  $P(\nu_\mu \rightarrow \nu_e)$  oscillation probability for a range of  $\delta_{CP}$  values. A circular modulation of the first oscillation peak (in both magnitude and position) is observed when varying throughout the allowable values of  $\delta_{CP}$ . The CP-conserving values of  $\delta_{CP} = 0, \pi$  have a lower(higher) oscillation maximum than the CP-violating values of  $\delta_{CP} = -\pi/2(\delta_{CP} = \pi/2)$ . A sub-dominant



**Figure 7.4.:** An illustration of the matter-induced effects on the oscillation probability, given as a function of neutrino energy and zenith angle. The top row of panels gives the  $P(\nu_e \rightarrow \nu_e)$  oscillation probability and the bottom row illustrates the  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  oscillation probability. The left column highlights the oscillation probability in a vacuum, whereas the middle and right column represents the oscillation probabilities when the four-layer fixed density PREM model is assumed. All oscillation probabilities assume the “Asimov A” set given in Table 2.2, but importantly, the right column sets an inverted mass hierarchy. The “matter resonance” effects at  $E_\nu \sim 5\text{GeV}$  can be seen in the  $P(\nu_e \rightarrow \nu_e)$  for normal mass hierarchy and  $P(\bar{\nu}_e \rightarrow \bar{\nu}_e)$  for inverted hierarchy.

<sup>2346</sup> shift in the energy of the oscillation peak is also present to aid in separating the two  
<sup>2347</sup> CP-conserving values of  $\delta_{CP}$ .

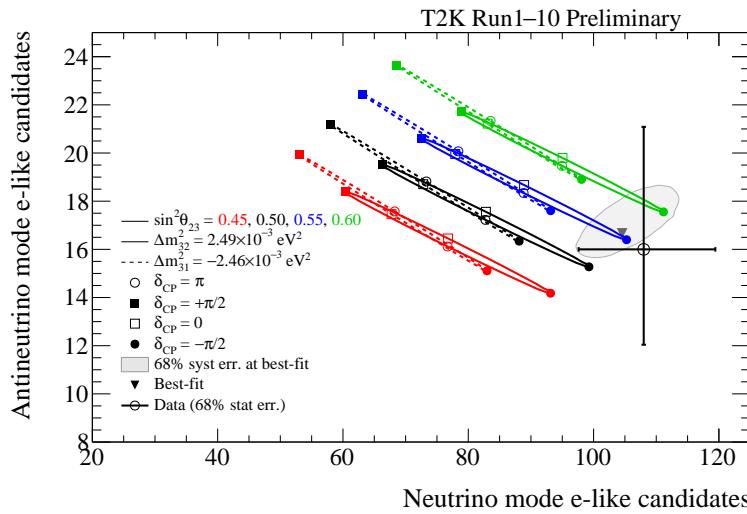
<sup>2348</sup> T2K’s sensitivity to the  $\sin^2(\theta_{23})$  and  $\Delta m_{23}^2$  is observed as a shape-based variation  
<sup>2349</sup> of the muon-like samples, as illustrated in Figure 7.5. The value of  $\Delta m_{32}^2$  laterally shifts  
<sup>2350</sup> the position of the oscillation dip (around  $E_\nu \sim 0.6\text{GeV}$ ) in the  $P(\nu_\mu \rightarrow \nu_\mu)$  oscillation



**Figure 7.5.:** The oscillation probability for beam neutrino events given as a function of neutrino energy. All oscillation parameters assume the “Asimov A” set given in Table 2.2 unless otherwise stated. Each panel represents a change in one of the oscillation parameters whilst keeping the remaining parameters fixed.

probability. A variation of  $\sin^2(\theta_{23})$  is predominantly observed as a vertical shift of the oscillation dip with second-order horizontal shifts being due to matter effects. The beam neutrinos have limited sensitivity to matter effects due to the relatively shorter baseline as well as the Earth’s mantle being a relatively low-density material (as compared to the Earth’s core). For some values of  $\delta_{CP}$ , the degeneracy in the number of e-like events allows the mass hierarchy to be resolved. This leads to a  $\delta_{CP}$ -dependent mass hierarchy sensitivity which can be seen in Figure 7.6.

Whilst all oscillation channels should be included for completeness, the computational resources required to run a fit are limited and any reasonable approximations



**Figure 7.6.:** The number of electron-like events in the FHC and RHC operating mode of the beam, as a function of the oscillation probabilities. Both normal hierarchy (Solid) and inverse hierarchy (Dashed) values of  $\Delta m_{23}^2$  are given.

which reduce the number of oscillation probability calculations that need to be made should be applied. The  $\nu_e \rightarrow \nu_{e,\mu,\tau}$  (and antineutrino equivalent) oscillations can be ignored for beam neutrinos as the  $\nu_e/\bar{\nu}_e$  fluxes are approximately two orders of magnitude smaller than the corresponding  $\nu_\mu/\bar{\nu}_\mu$  flux. Furthermore, as the peak neutrino energy of the beam is well below the threshold for charged current tau production ( $E_\nu = 3.5\text{GeV}$  [51], only a small proportion of the neutrinos produced in the beam have the required energy. For the few neutrinos that have sufficient energy, the oscillation probability is very small due to the short baseline. Whilst these approximations can be made for the beam neutrinos, the atmospheric flux of  $\nu_e$  is of the same order of magnitude as the  $\nu_\mu$  flux and the energy distribution of atmospheric neutrinos extends well above the tau production threshold.

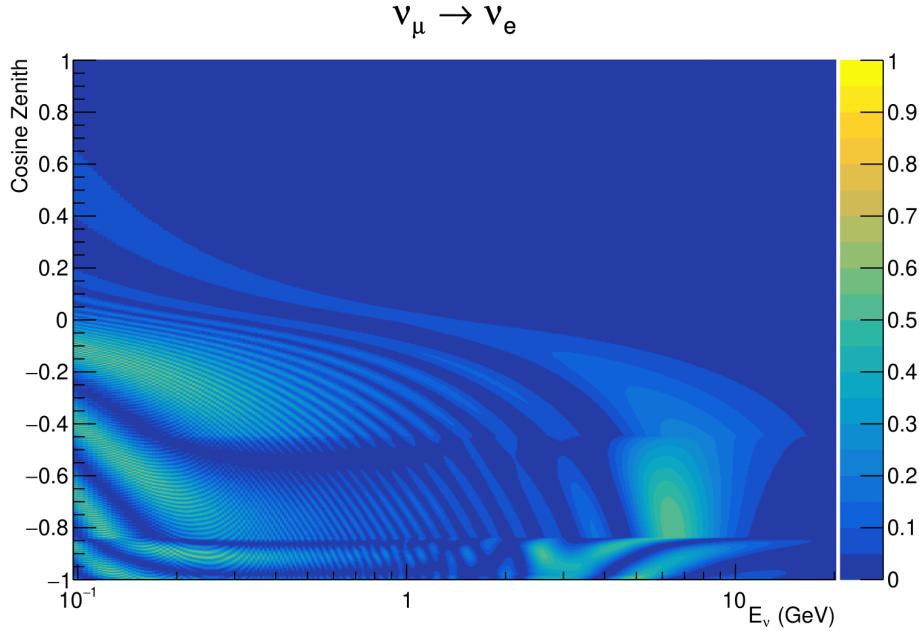
## 2371 7.2. Treatment of Fast Oscillations

2372 As shown in Figure 7.7, atmospheric neutrino oscillations have a significantly more  
2373 complex structure for upgoing neutrinos with energy below 1GeV. This is because the  
2374  $L/E$  dependence of the oscillation probability in this region induces rapid variations  
2375 for small changes in  $L$  or  $E$ . As discussed in section 7.1, this is also the region in which  
2376 atmospheric neutrinos have sensitivity to  $\delta_{CP}$ . In practice, the direction of the neutrino  
2377 is inferred from the direction of the final state particles traveling in the detector, which  
2378 can be poor for low-energy neutrino interactions. This creates a distinct difference  
2379 from the beam neutrinos where the position of the source is very precisely known.

2380 As a consequence of the unresolvable structure, an average oscillation probability  
2381 is observed in the subGeV upgoing region. This creates a computational problem; A  
2382 significantly large amount of Monte Carlo statistics would be required to accurately  
2383 predict the number of events if Monte Carlo averaging was the only technique used.  
2384 This section describes the ‘sub-sampling’ approach developed for this analysis and  
2385 compares it to the methodology used within the SK-only analysis.

2386 The official SK-only analysis uses the osc3++ oscillation parameter fitter [210].  
2387 To perform the fast oscillation averaging, it uses a ‘nearest-neighbour’ technique.  
2388 For a given neutrino event, the nearest twenty neighbours in reconstructed lepton  
2389 momentum and zenith angle are found and a distribution of their neutrino energies is  
2390 built. The RMS,  $\sigma$ , of this distribution is then used to compute an average oscillation  
2391 probability for the given neutrino Monte Carlo event.

2392 For the  $i^{th}$  event, the oscillation weight is calculated as



**Figure 7.7.:** The oscillation probability  $P(\nu_\mu \rightarrow \nu_e)$ , given as a function of neutrino energy and zenith angle, which highlights an example of the “fast” oscillations in the sub-GeV upgoing region.

$$W_i = \frac{1}{5}P(E_i, \bar{L}_i) + \frac{1}{5} \sum_{\beta=-1,-0.5,0.5,1} P(E_i + \beta\sigma_i, L_\beta), \quad (7.3)$$

where  $P(E, L)$  is the oscillation probability calculation for neutrino energy  $E$  and path length  $L$  and the two path lengths,  $\bar{L}_i$  and  $L_\beta$  are discussed below. All of the oscillation probability calculations are performed with a fixed zenith angle such that the same density profile is used.

The uncertainty in the production height is controlled by using an “average” production height,  $\bar{L}_i$ , which represents the average path length computed using twenty production heights taken from the Honda flux model’s prediction [45]. For a given event, the production heights are sampled in steps of 5% of their cumulative distribution function.  $L_\beta$  values are similarly calculated but instead use different combinations of four production heights,

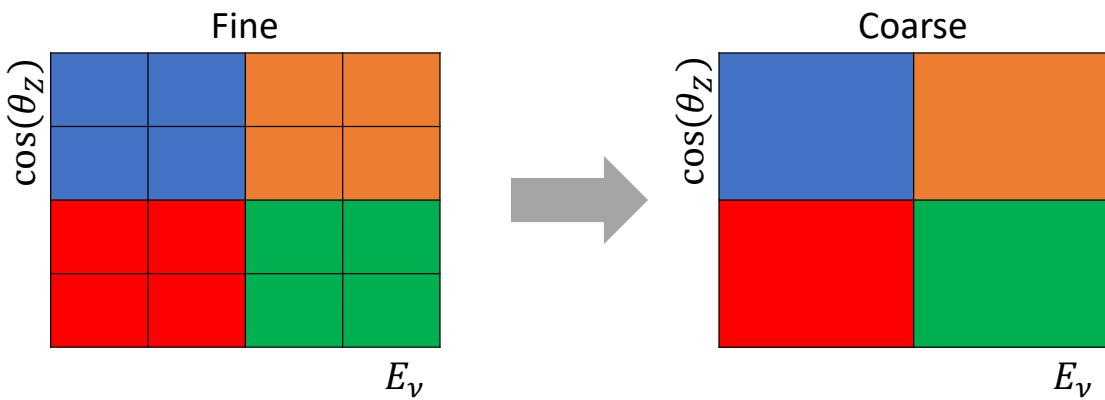
$$\begin{aligned}
 L_{-1.0} &= \frac{1}{4}L(45, 50, 55, 60), \\
 L_{-0.5} &= \frac{1}{4}L(35, 40, 65, 70), \\
 L_{+0.5} &= \frac{1}{4}L(25, 30, 75, 68), \\
 L_{+1.0} &= \frac{1}{4}L(15, 20, 85, 89).
 \end{aligned} \tag{7.4}$$

2403 This averaging technique works because of the inference between the zenith angle  
 2404 and the reconstructed direction of final state particles in the detector. For low-energy  
 2405 neutrinos, where the resolution of the true neutrino direction is poor,  $\sigma_i$  will be large,  
 2406 resulting in significant averaging effects. Contrary to this, the inferred direction of  
 2407 high-energy neutrinos will be much closer to the true value, meaning that  $\sigma_i$  will be  
 2408 smaller, culminating in small averaging effects.

2409 In practice, this technique is performed before the fit in order to deal with the  
 2410 computational cost. This is possible as the Osc3++ framework uses binned oscillation  
 2411 parameters rather than continuous so the oscillation parameters used in the fit are  
 2412 known prior to run-time. The framework used in this analysis uses continuous  
 2413 oscillation parameters, and due to the MCMC fitting technique, there is no way to  
 2414 know which oscillation parameter values will be selected *a priori*. Therefore, the  
 2415 oscillation parameter calculation has to be performed at run-time. Computing five  
 2416 oscillation probabilities per event would require far too many computational resources  
 2417 to be viable. Therefore SK technique can not be used within this analysis. However,  
 2418 the concept of the averaging technique can be taken from it.

2419 To perform a similar averaging as the SK analysis, a sub-sampling approach using  
 2420 binned oscillograms has been devised. The technique can be explained by considering  
 2421 a “fine” and “coarse” oscillogram. The fine oscillograms are used to define the array of

2422  $\cos(\theta_Z)$  and  $E_\nu$  used in the oscillation engine. The coarse oscillograms cover the same  
2423 phase-space but have fewer bins, where the value of a particular coarse bin is taken  
2424 as the linear average (flat prior in  $E_\nu$  and  $\cos(\theta_Z)$ ) of all fine bins which falls into it.  
2425 The coarse oscillogram is then used for determining the oscillation weight for a given  
2426 event. The binning which is used to calculate the oscillation probabilities, known as  
2427 the ‘fine’ binning, has  $N \times N$  subdivisions per coarse bin. Figure 7.8 illustrates the  
2428  $N = 2$  example where the assigned value to a coarse bin is the average of the four fine  
2429 bins which fall in that coarse bin. Whilst the coarse bin edges do not have to be linear  
2430 on either axis, the sub-division of the fine bins is linear over the range of a coarse bin.



**Figure 7.8.:** Illustration of the averaging procedure for  $N = 2$ . The oscillation probabilities calculated on the finer left binning are averaged to obtain the oscillation probabilities in the coarser right binning. These averaged oscillation probabilities with the coarser binning are then applied to each event during the fit.

2431 The coarse binning is defined with  $67 \times 52$  bins in true neutrino energy  $\times$  cosine  
2432 zenith. It is picked to be identical to that provided in [212]. In general, the binning is  
2433 logarithmically spaced in neutrino energy but has some hand-picked bin edges. Firstly,  
2434 the bin density around the matter resonance is smoothly increased around the matter  
2435 resonance region. This is to avoid smearing this region which can be well sampled by  
2436 the Monte Carlo. Secondly, bin edges are selected to hit  $0.4, 0.6, 1, 10, 30, 50, 100\text{GeV}$ .  
2437 This is to ensure that the Coulomb correction systematic and the atmospheric flux  
2438 systematics definitions in neutrino energy can be hit. The cosine zenith binning is

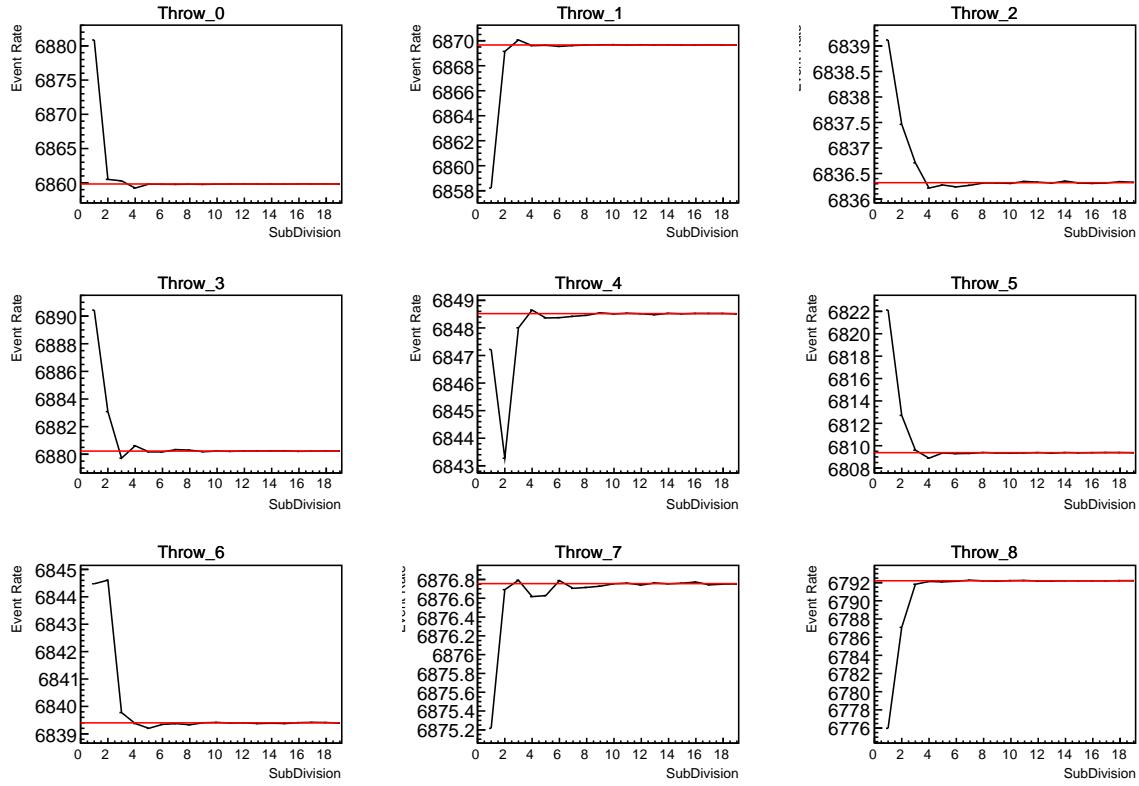
approximately linearly spaced across the allowable range but the values of layer transitions are hit precisely:  $-0.8376$  (core-mantle) and  $-0.4464$  (mantle/transition zone). Bins are spread further apart for downgoing events as this is a region unaffected by the fast oscillation wavelengths and reduces the total number of calculations required to perform the calculation.

The choice of  $N$  is justified based on two studies. Firstly, the variation of event rates of each sample is studied as a function of  $N$ . For a given set of oscillation parameters thrown from the PDG prior constraints (detailed in Table 2.1), the oscillation probabilities are calculated using a given value of  $N$ . Each sample is re-weighted and the event rate is stored. The value of  $N$  is scanned from 1, which corresponds to no averaging, to 19, which corresponds to the largest computationally viable subdivision binning. The event rate of each sample at large  $N$  is expected to converge to a stationary value due to the fine binning fully sampling the small-scale structure. Figure 7.9 illustrates this behaviour for the SubGeV\_elike\_0dcy sample for 9 different throws of the oscillation parameters.

Denoting the event rate for one sample for a given throw  $t$  at each  $N$  by  $\lambda_t^N$ , the average over all considered  $N$  values ( $\bar{\lambda}_t = \frac{1}{24} \sum_{N=1}^{24} \lambda_t^N$ ) is computed. The variance in the event rate at each  $N$  is then calculated as

$$\text{Var}[\lambda^N] = \frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t)^2 - \left[ \frac{1}{N_{\text{throws}}} \sum_{t=1}^{N_{\text{throws}}} (\lambda_t^N - \bar{\lambda}_t) \right]^2. \quad (7.5)$$

**The Below (and associated figures) has been added** In practice the following procedure is undertaken. For a particular throw, the difference between the event rate at a particular choice of  $N$  and the mean of the distribution is calculated. This is illustrated in Figure 7.10. This value is then calculated for all the 2000 throws,



**Figure 7.9.:** Event rate of the SubGeV\_elike\_0dcy sample as a function of the number of subdivisions per coarse bin. Each subplot represents the event rate of the sample at a different oscillation parameter set thrown from the PDG priors detailed in Table 2.1. The red line in each subplot represents the mean of the event rate over the different values of sub-divisions for that particular oscillation parameter throw.

generating a distribution of  $\lambda_t^N - \bar{\lambda}_t$ . This is repeated for each of the values of  $N$  considered within this study. The distributions of this value, for  $N = \{1, 5\}$ , are given in Figure 7.11. As expected, the distribution gets narrower and tends towards zero for the higher values of  $N$ .

#### 2465 End Review

2466 The aim of the study is to find the lowest value of  $N$  such that this variance is  
 2467 below 0.001. This utilises the width of the distributions given in Figure 7.11. This is  
 2468 the typical threshold used by T2K fitters to validate systematic implementation so has  
 2469 been set as the same criteria. The results of this study for each atmospheric sample  
 2470 used within this thesis are illustrated in Figure 7.12 for 2000 throws of the oscillation

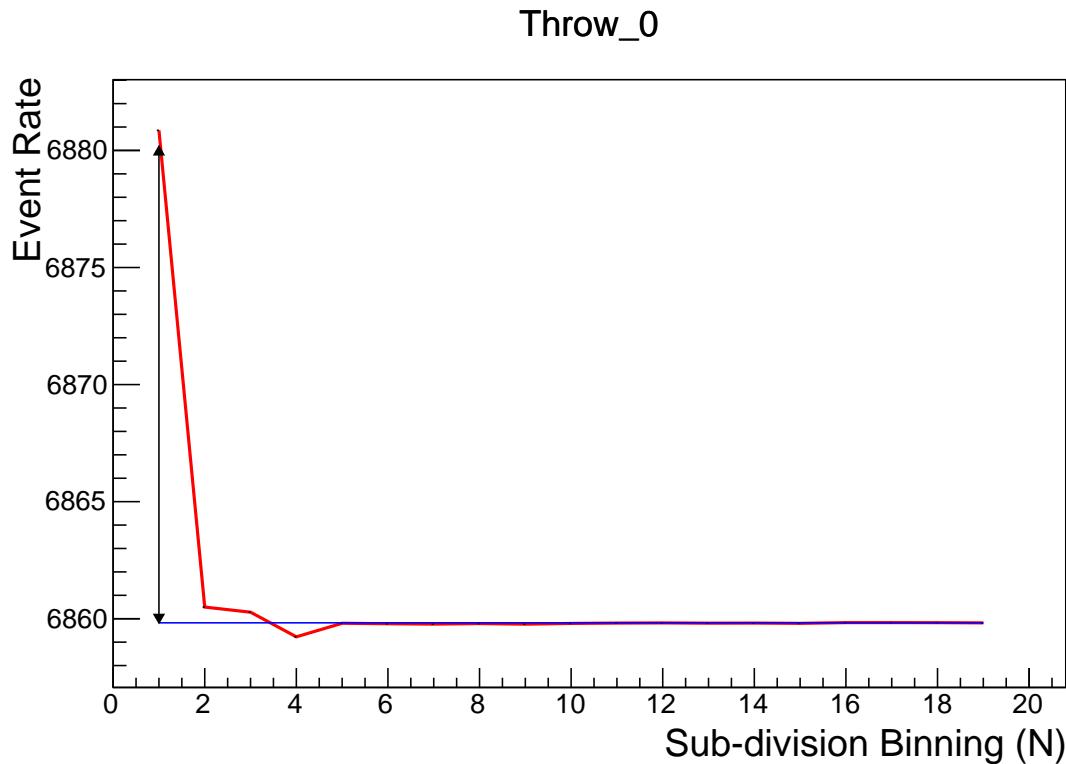
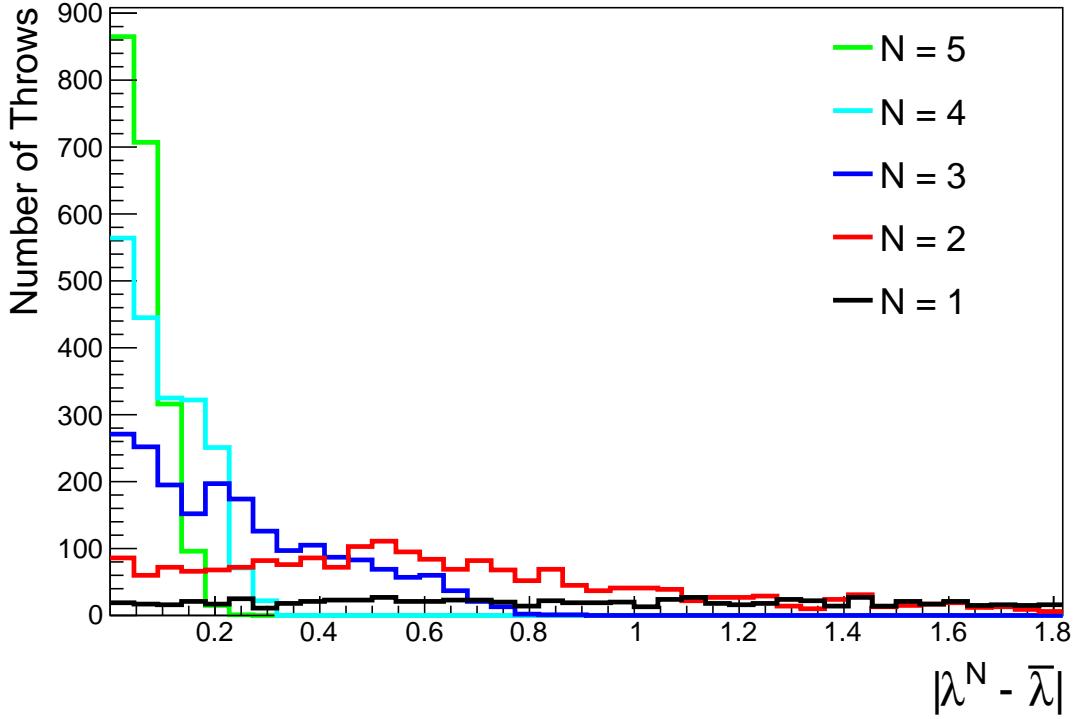


Figure 7.10.

parameters. As can be seen, the variance is below the threshold at  $N = 10$ , and is driven primarily by the SubGeV\_mulike\_1dcy and SubGeV\_elike\_0dcy samples.

The second study to determine the value of  $N$  is as follows. The likelihood for each sample is computed against an Asimov data set created with Asimov A oscillation parameters (Table 2.2). Following Equation 7.5, the variance of the log-likelihood over all considered  $N$  is computed. The results are shown in Figure 7.13.

A choice of  $N = 10$  sub-divisions per coarse bin has a variance in both event rate and log-likelihood residuals less than the required threshold of 0.001. The largest value of the likelihood variance is of order  $10^{-7}$ , corresponding to an error on the log-likelihood of about  $3 \times 10^{-4}$  which is small enough to be negligible for the oscillation analysis.



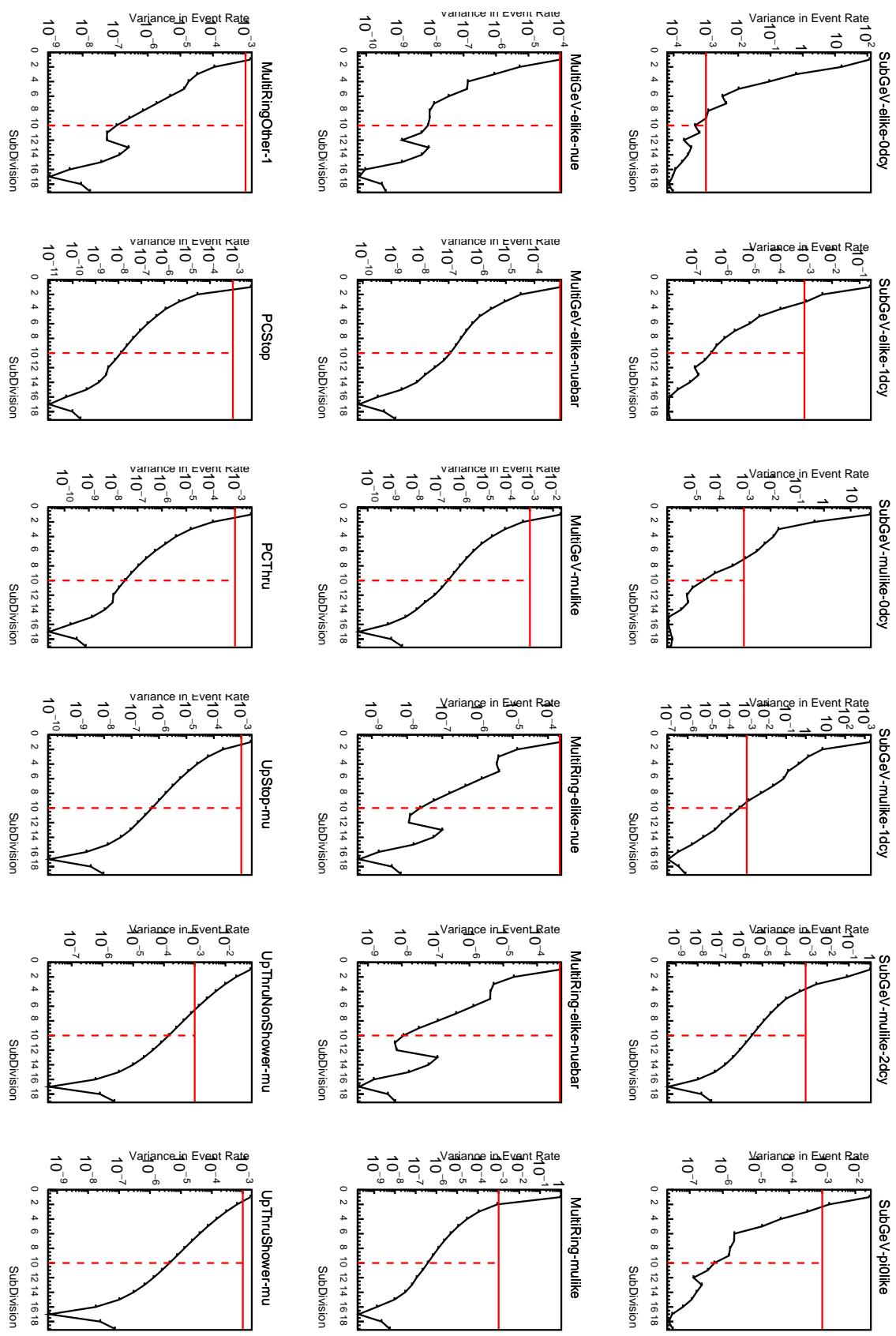
**Figure 7.11.**

2482     Figure 7.14 illustrates the effect of the smearing using  $N = 10$ . The fast oscillations  
 2483    in the sub-GeV upgoing region have been replaced with a normalisation effect whilst  
 2484    the large matter resonance structure remains.

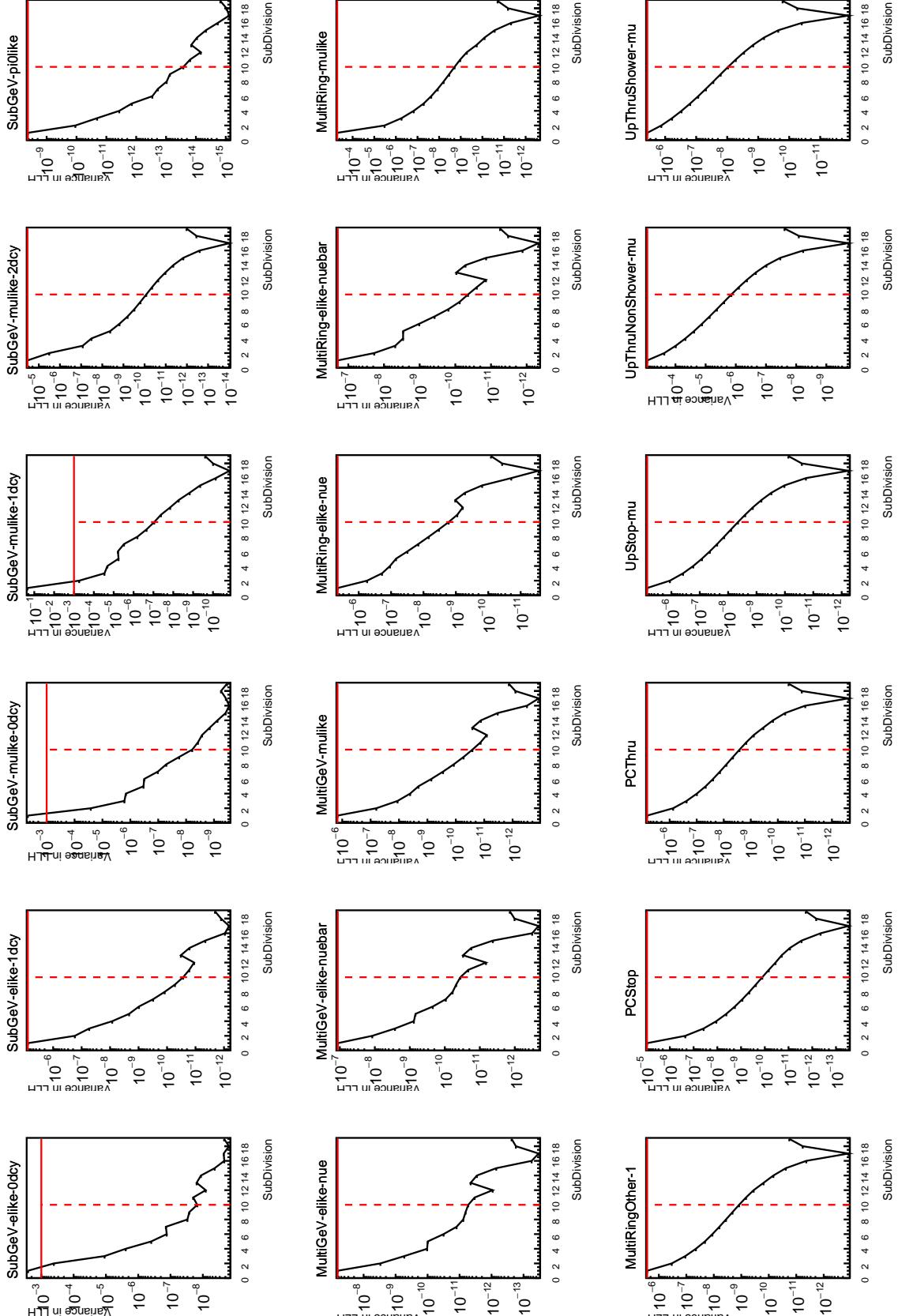
### 2485   **7.3. Calculation Engine**

2486   As previously discussed in section 7.2, the calculation of oscillation probabilities is per-  
 2487   formed at run-time due to utilising continuous oscillation parameters. Consequently,  
 2488   the time per calculation is crucial for fit performance. The initial fitting framework  
 2489   used for this analysis was developed with ProbGPU [213]. This is a GPU-only implemen-  
 2490   tation of the prob3 engine [214]. It is primarily designed for neutrino propagation in a  
 2491   beam experiment (single layer of constant density) with the atmospheric propagation  
 2492   code not being used prior to the analysis in this thesis.

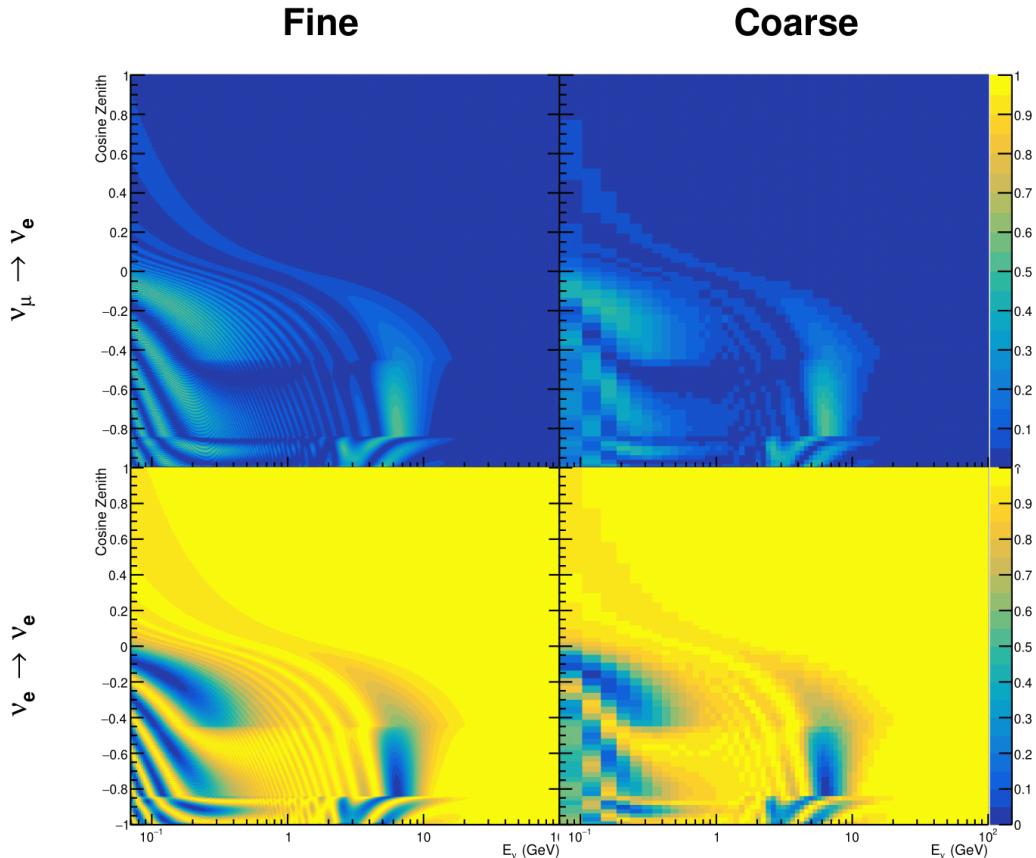
## Oscillation Probability Calculation



**Figure 7.12:** Variance of event rate for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line indicates the variance at a sub-division  $N = 10$ .



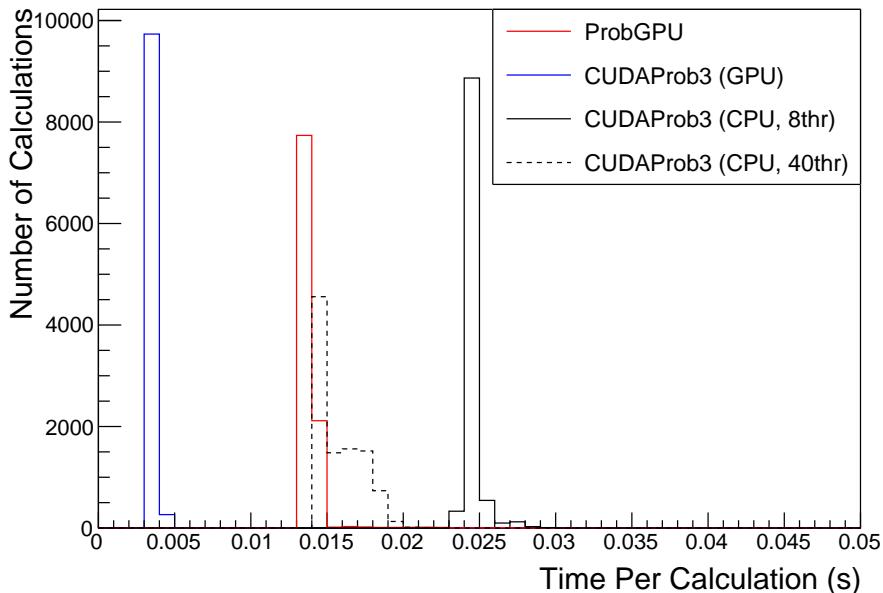
**Figure 7.13:** Variance of sample likelihood, when compared to 'Asimov data' set at Asimov A, for each atmospheric sample as a function of the number of sub-divisions per coarse bin. The solid red line indicates the 0.1% threshold and the dashed red line is a graphical indication of the variance at a sub-division  $N = 10$ .



**Figure 7.14.**: The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives the “fine” binning used to calculate the oscillation probabilities and the right column illustrates the “coarse” binning used to reweight the Monte Carlo events. The fine binning choice is given with  $N = 10$ , which was determined to be below the threshold from Figure 7.12 and Figure 7.13.

2493 Another engine, CUDAProb3 [215], has been implemented within the fitting frame-  
 2494 work used in this analysis. It has been specifically optimised for atmospheric neutrino  
 2495 oscillation calculation so does not contain the code to replace the beam oscillation  
 2496 calculation. The engine utilises object-orientated techniques as compared to the func-  
 2497 tional implementation of ProbGPU. This allows the energy and cosine zenith arrays to  
 2498 be kept on GPU memory, rather than having to load these arrays onto GPU memory  
 2499 for each calculation. General memory interfacing is one of the slowest tasks which

2500 GPUs can do, so being able to eliminate this significantly reduces the time required  
 2501 for calculation. This can be seen in Figure 7.15, where the GPU implementation of  
 2502 CUDAProb3 is approximately three times faster than the ProbGPU engine.



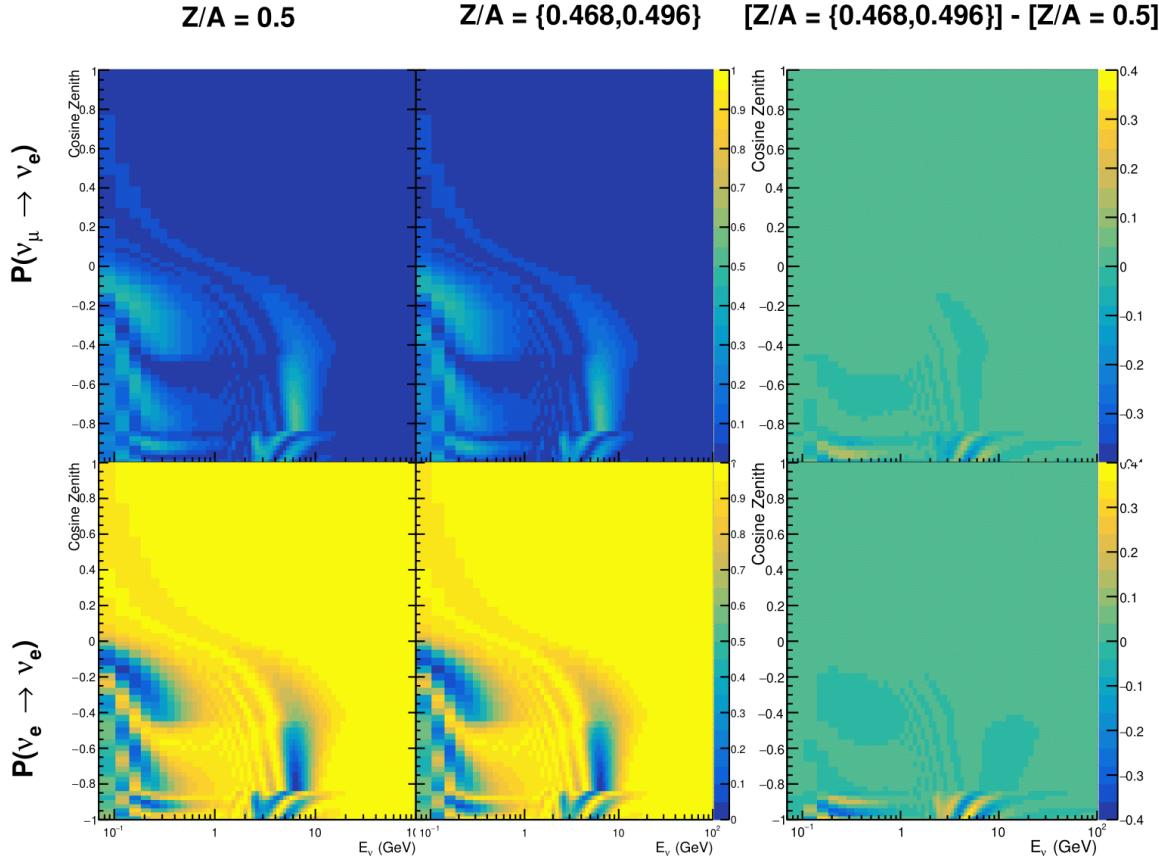
**Figure 7.15.:** The calculation time taken to both calculate the oscillation probabilities and fill the “coarse” oscilloscopes, following the technique given in section 7.2, for the CUDAProb3 and ProbGPU (Red) calculation engines. CUDAProb3 has both a GPU (Blue) and CPU (Black) implementation, where the CPU implementation is multithreaded. Therefore, 8-threads (solid) and 40-threads (dashed) configurations have been tested. Prob3, which is a CPU single-thread implementation has a mean step time of 1.142s.

2503 Another significant advantage of CUDAProb3 is that it contains a CPU multithreaded  
 2504 implementation which is not possible with the ProbGPU or prob3 engines. This elimi-  
 2505 nates the requirement for GPU resources when submitting jobs to batch systems. As  
 2506 illustrated in Figure 7.15, the calculation speed depends on the number of available  
 2507 threads. Using 8 threads (which is typical of the batch systems being used) is ap-  
 2508 proximately twice as slow as the ProbGPU engine implementation, but would allow  
 2509 the fitting framework to be run on many more resources. This fact is utilised for any  
 2510 SK-only fits but GPU resources are required for any fits which include beam samples  
 2511 due to the ProbGPU requirement. Based on the benefits shown by the implementation

2512 in this section, efforts are being placed into including linear propagation for beam  
2513 neutrino propagation into the engine [216].

## 2514 7.4. Matter Density Profile

2515 For an experiment observing atmospheric neutrinos propagating through the Earth, a  
2516 model of the Earth's density profile is required. The model used within this analysis is  
2517 the Preliminary Reference Earth Model (PREM) [211], as illustrated in Figure 7.1. As  
2518 discussed in section 7.1, the propagator used within the calculation engine requires  
2519 constant density layers. To follow the official SK-only analysis [210], the average  
2520 density of each layer has been taken from the PREM model. Table 7.1 documents  
2521 the density and radii of the layers used within this approximation. The density  
2522 measurements provided in the PREM model are provided in terms of mass density,  
2523 whereas neutrino oscillations are sensitive to the electron number density. This value  
2524 can be computed as the product of the chemical composition, or the  $Z/A$  value, and  
2525 the mass density of each layer. Currently, the only way to calculate the chemical  
2526 composition value for layers close to the Earth's core is through neutrino oscillations.  
2527 The chemical composition of the upper layers of the Earth's Mantle and the Transition  
2528 zone is well known due to it being predominantly pyrolite which has a chemical  
2529 composition value of 0.496 [217]. The components of the Earth's core region are less  
2530 well known. Consequently, the chemical composition dial for the core layers is set to a  
2531 value of 0.468, as calculated in [218]. This value is assigned a Gaussian error with a  
2532 standard deviation equivalent to the difference in chemical composition in core and  
2533 mantle layers. Figure 7.16 illustrates the effect of moving from the  $Z/A = 0.5$  method  
2534 which is used in the official SK-only analysis [210] to these more precise values.



**Figure 7.16.:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_e)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the constant  $Z/A = 0.5$  approximation which is used in the official SK-only analysis. The middle column gives the probabilities where  $Z/A = [0.468, 0.498]$  values are used, as given in Table 7.1. The right column illustrates the difference in oscillation probability between the two different techniques.

2535     The beam oscillation probability in this thesis uses a baseline of 295km, density  
 2536      $2.6\text{g/cm}^3$ , and chemical composition 0.5 as is done by the official T2K-only analysis  
 2537     [219].

2538     Whilst the propagator requires a fixed density layer model of the Earth, the density  
 2539     only has to be fixed for a specific  $E_\nu \times \cos(\theta_Z)$  bin in a given layer. As the density is a  
 2540     function of radius, which is a function of the direction in which a neutrino propagates,

2541 a better approximation of the PREM model can be made if a  $\cos(\theta_Z)$ -specific density is  
2542 calculated.

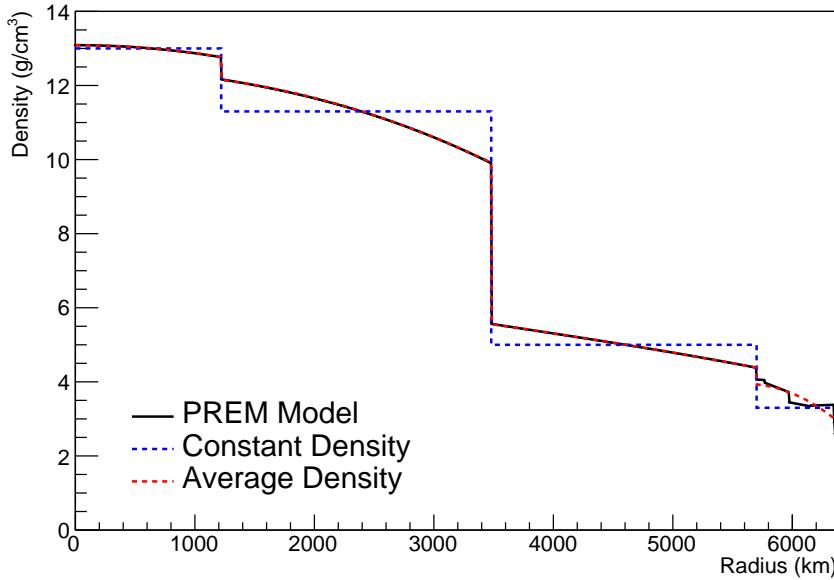
2543 To achieve this, the average density,  $\langle \rho \rangle_i$ , in the  $i^{th}$  layer, is calculated as the density,  
2544  $\rho(t)$ , integrated over the track a given  $\cos(\theta_Z)$ ,

$$\langle \rho \rangle_i = \frac{1}{t_{i+1} - t_i} \int_{t_i}^{t_{i+1}} \rho(t) dt \quad (7.6)$$

2545 where  $t_i$  are the intersection points between each layer and  $t$  is the path length of  
2546 the trajectory across the layer.

2547 The oscillation probability calculation speed is approximately linear in the number  
2548 of layers invoked within the Earth model. Therefore a four-layer model is still utilized  
2549 with the only difference to the official SK-only analysis being that the four-layer model  
2550 used for each value of  $\cos(\theta_Z)$  is different. Following the method outlined in [220],  
2551 a four-layer piecewise quadratic polynomial is fit to the PREM model for the four  
2552 layers defined in Table 7.1. This fit was not performed by the author of the thesis  
2553 and is documented in [212]. The coefficients of the quadratic fit to each layer are  
2554 given in Table 7.2 with the final distribution illustrated in Figure 7.17. The quadratic  
2555 approximation is clearly much closer to the PREM model as compared to the constant  
2556 density approximation.

2557 The effect of using the quadratic density per  $\cos(\theta_Z)$  model is highlighted in  
2558 Figure 7.18. The slight discontinuity in the oscillation probability around  $\cos(\theta_Z) \sim -$   
2559 0.45 in the fixed density model, which is due to the transition to mantle layer boundary,  
2560 has been reduced. This is expected as the difference in the density across this boundary  
2561 is significantly smaller in the quadratic density model as compared to the constant  
2562 density model. Whilst the difference in density across the other layer transitions

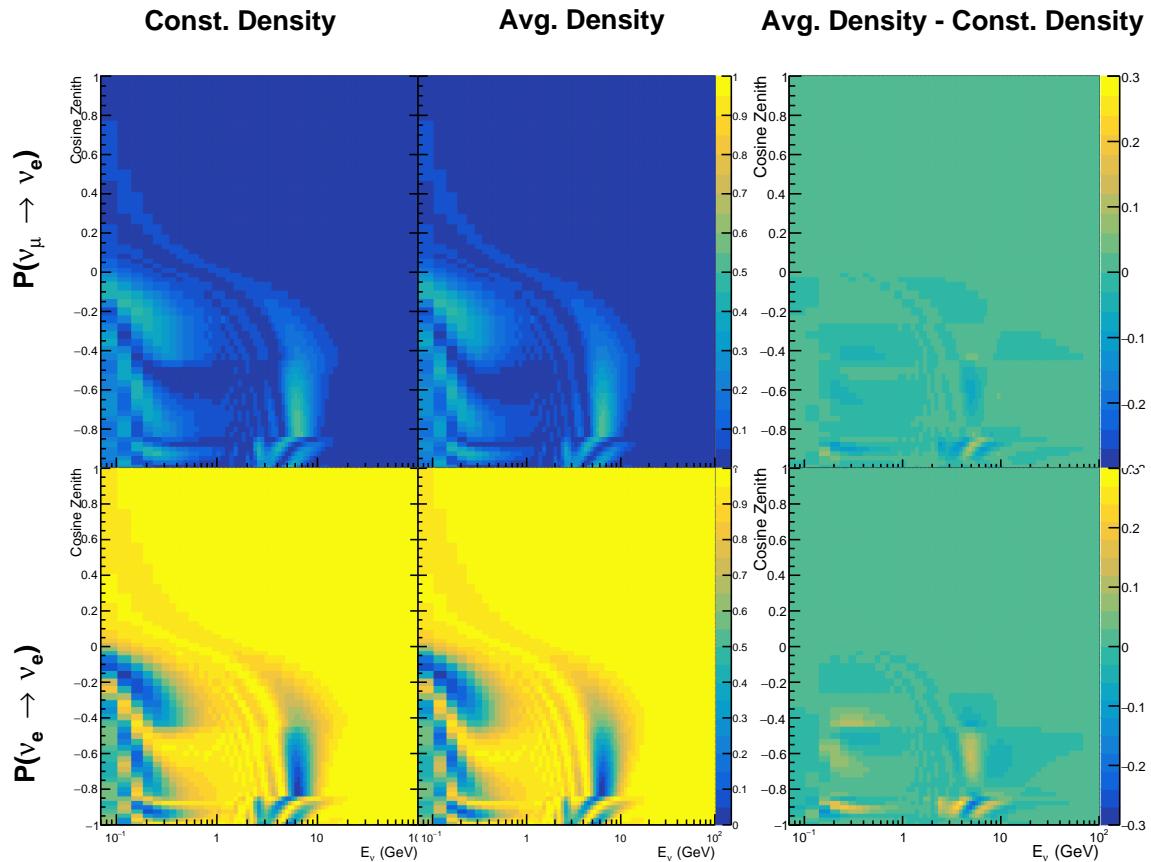


**Figure 7.17.:** The density of the Earth given as a function of the radius, as given by the PREM model (Black), the constant density four-layer approximation (Blue), as used in the official SK-only analysis, and the quadratic approximation of the PREM model (Red).

Layer	Outer Radius [km]	Density [g/cm <sup>3</sup> ]
Inner Core	1220	$13.09 - 8.84x^2$
Outer Core	3480	$12.31 + 1.09x - 10.02x^2$
Lower Mantle	5701	$6.78 - 1.56x - 1.25x^2$
Transition Zone	6371	$-50.42 + 123.33x - 69.95x^2$

**Table 7.2.:** The quadratic polynomial fits to the PREM model for four assumed layers of the PREM model. The fit to calculate the coefficients is given in [212], where  $x = R/R_{Earth}$ .

is reduced, there is still a significant difference. This means the discontinuities in the oscillation probabilities remain but are significantly reduced. However, as the quadratic density approximation matches the PREM model well in this region, these discontinuities are due to the Earth model rather than an artifact of the oscillation calculation.



**Figure 7.18.:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_\mu)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where the four-layer constant density approximation is used. The middle column gives the probabilities where the density is integrated over the trajectory, using the quadratic PREM approximation, for each  $\cos(\theta_Z)$  is used. The right column illustrates the difference in oscillation probability between the two different techniques.

## 2568 7.5. Production Height Averaging

2569 As discussed in section 7.1, the height at which the cosmic ray flux interacts in the  
2570 atmosphere is not known on an event-by-event basis. The production height can vary  
2571 from the Earth’s surface to  $\sim 50\text{km}$  above that. The SK-only analysis methodology  
2572 (described in section 7.2) for including the uncertainty on the production height is  
2573 to include variations from the Honda model when pre-calculating the oscillation  
2574 probabilities prior to the fit. This technique is not possible for this analysis which  
2575 uses continuous oscillation parameters that can not be known prior to the fit. Conse-  
2576 quently, an analytical averaging technique was developed in [212]. The author of this  
2577 thesis was not responsible for the derivation of the technique but has performed the  
2578 implementation and validation of the technique for this analysis alone.

2579 Using the 20 production heights per Monte Carlo neutrino event, provided as 5%  
2580 percentiles from the Honda flux model, a production height distribution  $p_j(h|E_\nu, \cos \theta_Z)$   
2581 is built for each neutrino flavour  $j = \nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu$ . In practice, a histogram is filled with  
2582 20 evenly spaced bins in production height  $h$  between 0 and 50km. The neutrino energy  
2583 and cosine zenith binning of the histogram is the same as that provided in section 7.2.  
2584 The average production height,  $\bar{h} = \int dh \frac{1}{4} \sum_j p_j(h|E_\nu, \cos(\theta_Z))$ , is calculated. The  
2585 production height binning of this histogram is then translated into  $\delta t(h) = t(\bar{h}) - t(h)$ ,  
2586 where  $t(h)$  is the distance travelled along the trajectory.

2587 For the  $i^{\text{th}}$  traversed layer, the transition amplitude,  $D_i(t_{i+1}, t_i)$ , is computed. The  
2588 time-ordered product of these is then used as the overall transition amplitude via

$$A(t_{n+1}, t_0) = D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1) D_0(t_1, t_0), \quad (7.7)$$

2589

where,

$$\begin{aligned} D_n(t_{n+1}, t_n) &= \exp[-iH_n(t_{n+1} - t_n)] \\ &= \sum_{k=1}^3 C_k \exp[ia_k(t_{n+1} - t_n)] \end{aligned} \quad (7.8)$$

2590

is expressed as a diagonalised time-dependent solution to the Schrodinger equation.

2591

The  $0^{th}$  layer is the propagation through the atmosphere and is the only term that depends on the production height. Using the substitution  $t_0 = t(\bar{h}) - \delta t(h)$ , it can be shown that

2592

Thus Equation 7.7 becomes

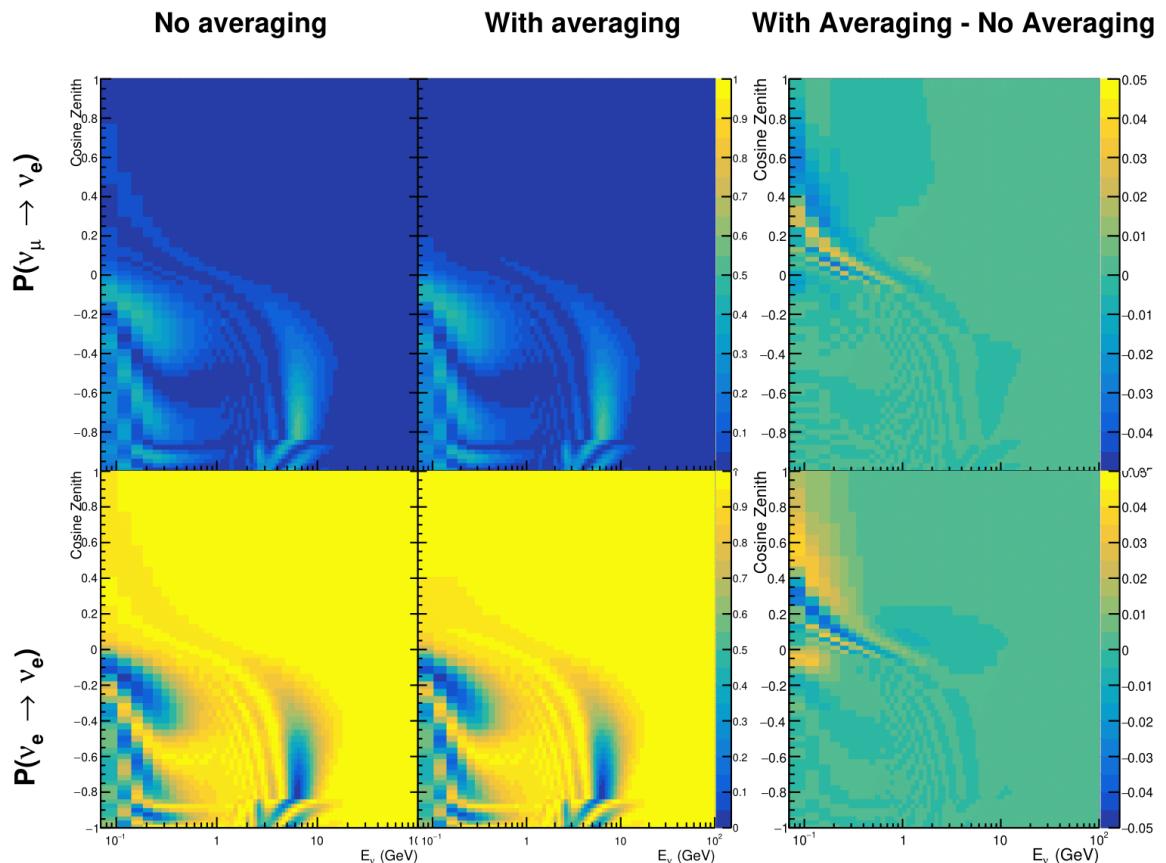
$$\begin{aligned} A(t_{n+1}, t_0) &= D_n(t_{n+1}, t_n) \dots D_1(t_2, t_1) D_0(t_1, \bar{h}) D(\delta t) \\ &= A(t_{n+1}, \bar{h}) \sum_{k=1}^3 C_k \exp[ia_k \delta t], \\ &= \sum_{k=1}^3 B_k \exp[ia_k \delta t]. \end{aligned} \quad (7.10)$$

2593

The oscillation probability averaged over production height is then calculated as

$$\begin{aligned}
 \bar{P}(\nu_j \rightarrow \nu_i) &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) P(\nu_j \rightarrow \nu_i) \\
 &= \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) A(t_{n+1}, t_0) A^*(t_{n+1}, t_0) \\
 &= \sum_{km} (B_k)_{ij} (B_m)_{ij}^* \int d(\delta t) p_j(\delta t | E_\nu, \cos \theta_Z) \exp[i(a_k - a_m)\delta t]
 \end{aligned} \tag{7.11}$$

2596 In practice, implementation in CUDAProb3 [215] is relatively straightforward as  
 2597 the majority of these terms are already calculated in the standard oscillation calculation.  
 2598 Figure 7.19 illustrates the results of the production height averaging. As expected,  
 2599 the main effect is observed in the low-energy downward-going and horizontal-going  
 2600 events. Upward-going events have to travel the radius of the Earth,  $R_E = 6371\text{km}$ ,  
 2601 where the production height uncertainty is a small fraction of the total path length.



**Figure 7.19.:** The oscillation probability,  $P(\nu_\mu \rightarrow \nu_e)$  (top row) and  $P(\nu_e \rightarrow \nu_\mu)$  (bottom row), given as a function of neutrino energy and zenith angle. The left column gives probabilities where a fixed production height of 25km is used. The middle column gives the probabilities where the production height is analytically averaged. The right column illustrates the difference in oscillation probability between the two different techniques.

2602 **Chapter 8**

2603 **Oscillation Analysis**

2604 **8.1. Likelihood Calculation**

2605 This analysis performs a joint oscillation parameter fit of the ND280, and the SK  
2606 atmospheric samples.

2607 Once the Monte Carlo predictions of each beam and atmospheric sample has been  
2608 built, following from chapter 6, a likelihood needs to be constructed. This is done  
2609 by comparing the Monte Carlo prediction to “data”. The data can consist of either  
2610 an Asmiov Monte Carlo prediction, which is typically used for sensitivity studies,  
2611 or real data. The Monte Carlo prediction is calculated at a particular point,  $\vec{\theta}$ , in the  
2612 model parameter space,  $N_i^{MC} = N_i^{MC}(\vec{\theta})$ . Both the data and Monte Carlo spectra are  
2613 binned, where the  $i^{th}$  bin content is represented by  $N_i^D$  and  $N_i^{MC}$ , respectively. The bin  
2614 contents for the beam near detector, beam far detector and atmospheric samples are  
2615 denoted with  $ND$ ,  $FD$  and  $Atm$ , respectively. The binning index,  $i$ , runs over all the  
2616 bins within the sample and all samples with that set. Taking the beam far detector  
2617 samples as example, it would run over all the reconstructed neutrino energy bins in all  
2618 samples (FHC1R $\mu$ , RHC1R $\mu$ , etc.). The likelihood calculation between data and Monte  
2619 Carlo for a particular bin follows a Poisson distribution, where the data is treated as a  
2620 fluctuation of the simulation.

2621 Following the T2K analysis presented in [80], the likelihood contribution from the  
2622 near detector also includes a Monte Carlo statistical uncertainty term, derived from

the Barlow and Beeston statistical treatment [221, 222]. In addition to treating the data as a fluctuation of the Monte Carlo prediction, it includes a contribution from the likelihood that the generated simulation is a statistical fluctuation of the actual true simulation assuming infinite statistics. The technical implementation of this additional likelihood term is documented in [195]. The term is defined as,

$$\frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2}, \quad (8.1)$$

where  $\beta_i$  represents a scaling parameter for each bin  $i$ , which is a value based on the amount of Monte Carlo statistics in a bin [195].  $\sigma_{\beta_i} = \sqrt{\sum_i w_i^2 / N_i^{MC}}$ , and  $\sqrt{\sum_i w_i^2}$  represents the sum of the square of the weights of the Monte Carlo events which fall into bin  $i$ .

Additional contributions to the likelihood come from the variation of the systematic model parameters. For those parameters with well-motivated uncertainty estimates, a covariance matrix,  $V$  describes the prior knowledge of each parameter as well as any correlations between the parameters. Due to the technical implementation, a single covariance matrix describes each “block” of model parameters, e.g. beam flux systematics. For simplicity, the covariance matrix associated with the  $k^{th}$  block is denoted  $V^k$ . This substitution results in  $\vec{\theta} = \sum_k^{N_b} \vec{\theta}^k$  and  $V = \sum_k^{N_b} V^k$ , for  $N_b$  number of blocks describing: oscillation parameters, beam flux, atmospheric flux, neutrino interaction, near detector, beam far detector and atmospheric far detector systematics detailed in section 6.4. The number of parameters in the  $k^{th}$  block is defined as  $n(k)$ .

The final likelihood term is defined as,

$$\begin{aligned}
& -\ln(\mathcal{L}) = & (8.2) \\
& \sum_i^{\text{NDbins}} N_i^{\text{ND},MC}(\vec{\theta}) - N_i^{\text{ND},D} + N_i^{\text{ND},D} \times \ln \left[ N_i^{\text{ND},D} / N_i^{\text{ND},MC}(\vec{\theta}) \right] + \frac{(\beta_i - 1)^2}{2\sigma_{\beta_i}^2} \\
& + \sum_i^{\text{FDbins}} N_i^{\text{FD},MC}(\vec{\theta}) - N_i^{\text{FD},D} + N_i^{\text{FD},D} \times \ln \left[ N_i^{\text{FD},D} / N_i^{\text{FD},MC}(\vec{\theta}) \right] \\
& + \sum_i^{\text{Atmbins}} N_i^{\text{Atm},MC}(\vec{\theta}) - N_i^{\text{Atm},D} + N_i^{\text{Atm},D} \times \ln \left[ N_i^{\text{Atm},D} / N_i^{\text{Atm},MC}(\vec{\theta}) \right] \\
& + \frac{1}{2} \sum_k^{N_b} \sum_i^{n(k)} \sum_j^{n(k)} (\vec{\theta}^k)_i (V^k)_{ij}^{-1} (\vec{\theta}^k)_j.
\end{aligned}$$

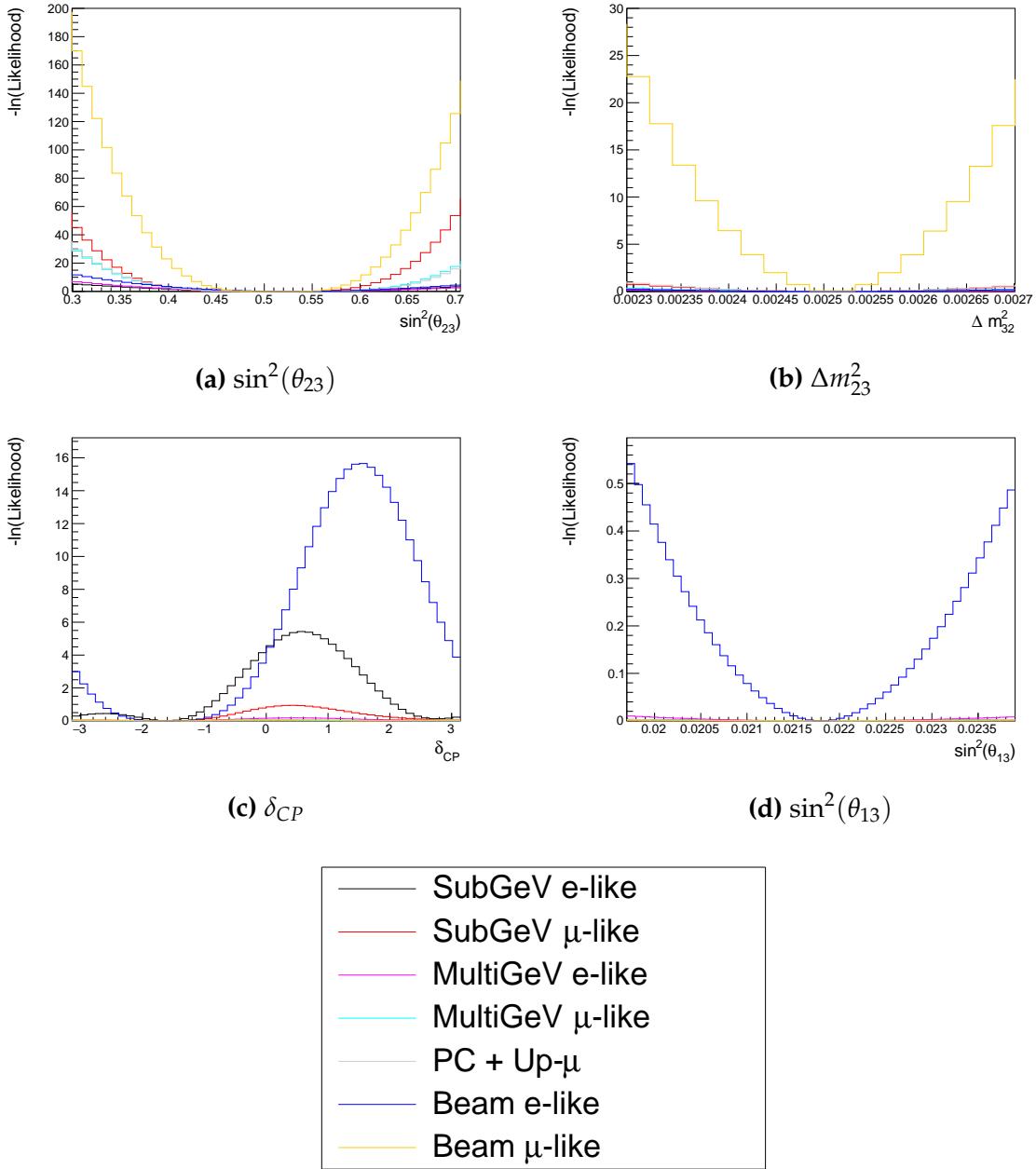
2643 This is the value determined at each step of the MCMC to build the posterior  
2644 distribution, as discussed in chapter 4.

**2645 8.1.1. Likelihood Scans**

2646 Using the defintion of the likelihood presented in section 8.1, the response of each  
 2647 sample to a variation particular parameter can be studied. Figure 8.1 presents the  
 2648 variation of all the samples (beam and atmospheric) at SK. Each plot represents a  
 2649 “scan”, where a particular parameter is scanned in some range. The “data” being  
 2650 used within the definition of the likelihood equation is built using the Asimov A  
 2651 oscillation parameter values defined in Table 2.2 alongside the pre-fit dial values as  
 2652 discussed in subsection 6.4.3. Due to the correlations between oscillation parameters,  
 2653 the value of  $\chi^2 \sim 1$  does not equate to the typical  $1\sigma$  sensitivity. However, it does give  
 2654 an indication of which samples response the strongest to a variation in the oscillation  
 2655 parameters. The point at which the likelihood tends to zero illustrates the value of the  
 2656 parameter used to build the Asimov data prediction. The likelihood scans only include  
 2657 the sample response and ignore the penalty contribution term from the variation of  
 2658 the parameter.

2659 The response to  $\Delta m_{23}^2$  is much larger in beam samples, specifically  $\mu$ -like samples,  
 2660 compared to atmospheric samples. This is to be expected as the beam neutrino  
 2661 energy can be specifically tuned to match the maximal disappearance probability.  
 2662 As discussed in section 7.1, the determination of the mass hierarchy is significantly  
 2663 enhanced when using the atmospheric samples due to them transitioning through the  
 2664 Earth’s core. So whilst the atmospheric samples do not add much information to the  
 2665 constraint of  $|\Delta m_{32}^2|$  beyond that of the beam analysis, they do enhance the ability to  
 2666 determine the sign of the parameter.

2667 The sensitivity to  $\sin^2(\theta_{23})$  is again dominated by the T2K experiment. However,  
 2668 the difference in the response for atmospheric and beam samples is much smaller. Con-  
 2669 sequently, one would expect that the joint fit would become more sensivity to  $\sin^2(\theta_{23})$



**Figure 8.1.:** The response of the likelihood, as defined in section 8.1, illustrating the response of the samples to the oscillation parameters.  $\Delta m_{12}^2$  and  $\sin^2(\theta_{12})$  are negated because these samples have no sensitivity to those parameters. The Asimov data set is built using the pre-fit dial values assuming Asimov A oscillation parameters defined in Table 2.2. DB: Need finer binning on delmsq23

than just T2K experiment alone. The summed response over all atmospheric samples becomes comparable to that of the muon-like beam samples. For this particular choice of Asimov point, the only samples which respond to the  $\sin^2(\theta_{13})$  parameter are the

2673 electron-like beam samples. Consequently, no increase in sensitivity beyond that of  
 2674 the T2K-only analysis is expected at that Asimov point. The  $\Delta m_{12}^2$  and  $\sin^2(\theta_{12})$  are  
 2675 not considered as there is simply no sensitivity in any sample considered within this  
 2676 analysis.

2677 As discussed, the correlations between oscillation parameters induce marginali-  
 2678 sation effects within the response of the likelihood. That is to say, the response to  
 2679  $\delta_{CP}$  is affected by the choice of  $\sin^2(\theta_{13})$  or  $\sin^2(\theta_{23})$ . The two-dimensional scans of  
 2680 the appearance ( $\sin^2(\theta_{13}) - \delta_{CP}$ ) and disappearance ( $\sin^2(\theta_{23}) - \Delta m_{23}^2$ ) parameters are  
 2681 illustrated in Figure 8.2 and Figure 8.3, respectively.

2682 The appearance log-likelihood scans show the distinct difference in how the beam  
 2683 and atmospheric samples respond. The beam samples have an approximately constant  
 2684 width of the  $2\sigma$  and  $3\sigma$  contours, throughout all ranges of  $\delta_{CP}$ . The atmospheric  
 2685 samples response to  $\delta_{CP}$  is very strongly correlated to the choice of  $\sin^2(\theta_{13})$ , with  
 2686 the strongest constraints around  $\delta_{CP} \sim 1$ . Consequently, this difference allows some  
 2687 of the degeneracy in a beam-only fit to be broken. Comparing the beam and joint  
 2688 fit log-likelihood scans, the  $2\sigma$  continuous contour in  $\delta_{CP}$  for beam samples is broken  
 2689 when the atmospheric samples are added. Furthermore, the width of the  $3\sigma$  contours  
 2690 also becomes dependent upon the value of  $\delta_{CP}$ . Whilst these are encouraging results  
 2691 for the joint fit, these are not sensitivity measurements as the nuisance parameters are  
 2692 fixed.

2693 The disappearance log-likelihood scans in  $\sin^2(\theta_{23}) - \Delta m_{23}^2$  space show the expected  
 2694 result when considering the one-dimensional scans already discussed. The uncertainty  
 2695 on the width of  $|\Delta m_{32}^2|$  is mostly driven by the beam-only sensitivities. However, the  
 2696 width of this contour in the inverted mass region ( $\Delta m_{32}^2 < 0$ ) is significantly reduced  
 2697 due to the ability of the atmospheric samples to select the correct mass hierarchy (these  
 2698 log-likelihood scans use the Asimov A oscillation probabilities which assumes true

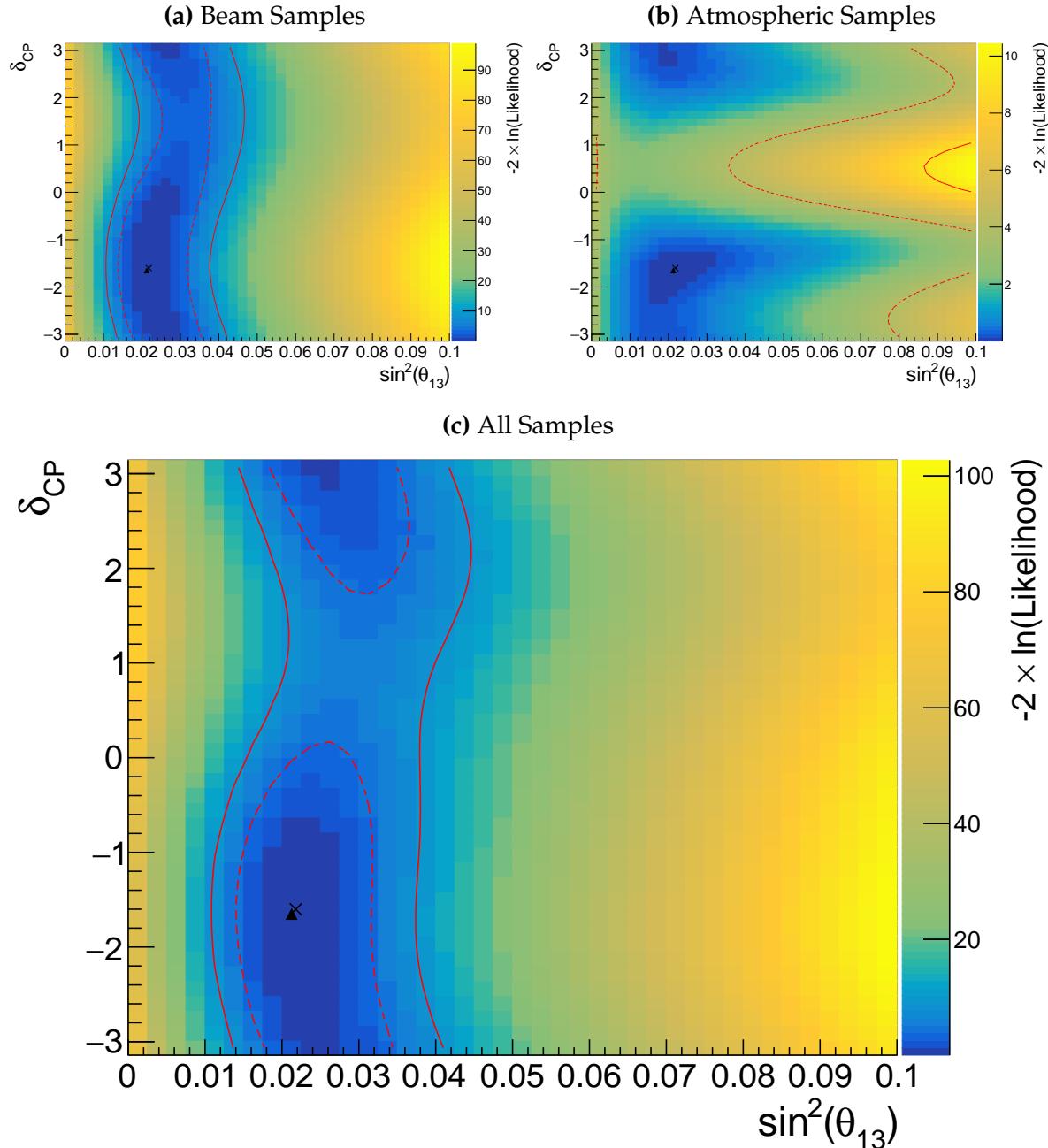
2699 normal hierarchy). The width of the uncertainty in  $\sin^2(\theta_{23})$  is also reduce compared  
2700 to a beam-only analysis, with a further decrease in the inverted hierarchy region due  
2701 to mass hierarchy determination.

2702 In addition to the oscillation parameters, the response to the systematic model  
2703 parameters can also be considered. Due to the correlated cross section model, the most  
2704 informative DB: Finish this

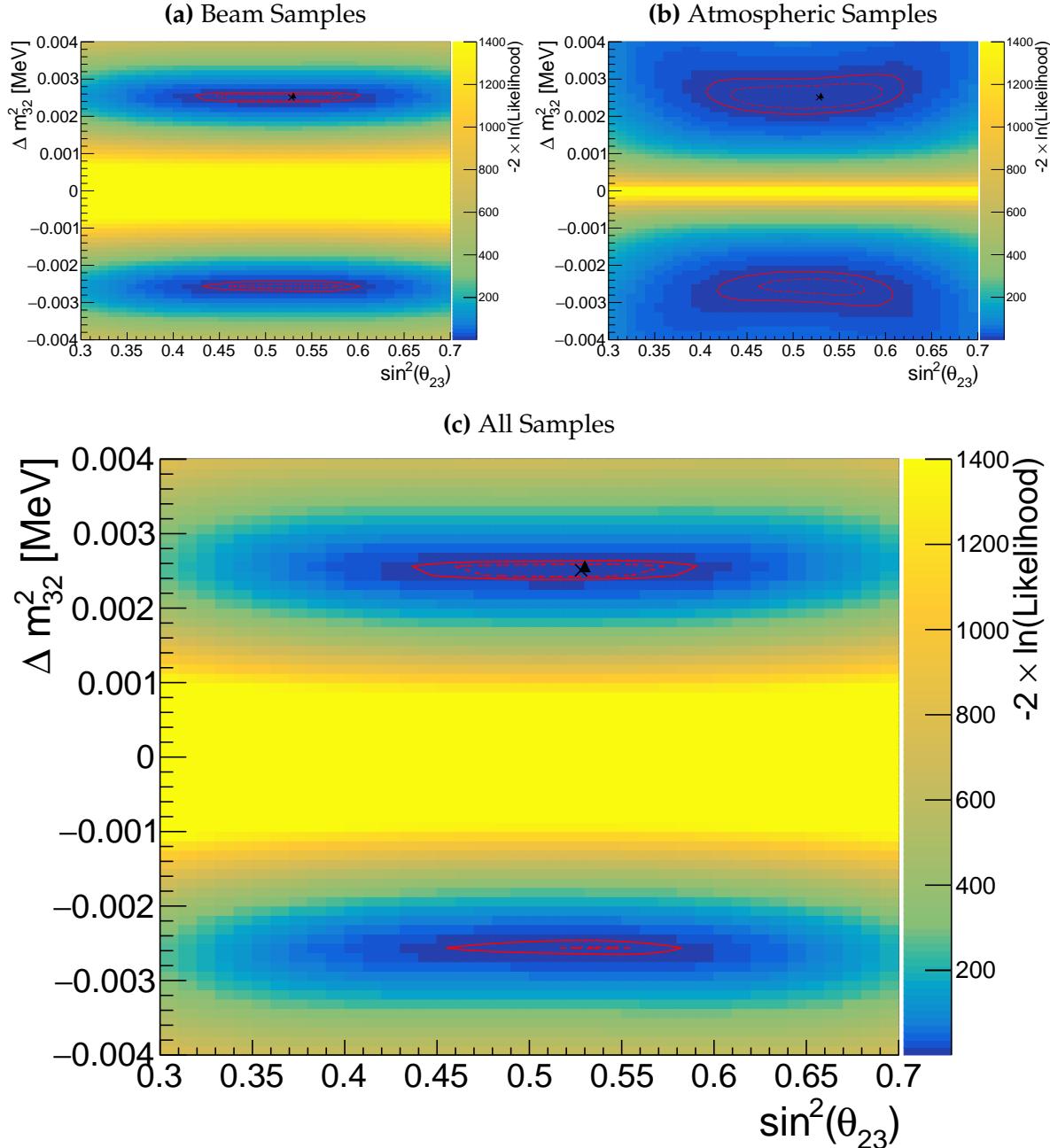
## 2705 8.2. Monte Carlo Prediction

2706 Using the three sets of dial values defined in subsection 6.4.3, the predicted event rates  
2707 for each sample are defined in Table 8.1. Both the oscillated event rates assuming  
2708 Asimov A oscillation parameters (defined in Table 2.2) and the un-oscillated event  
2709 rates are given.

2710 Generally, samples which target CCQE interaction modes observe a decrease in  
2711 prediction when using the pre-fit dial values. This is in accordance with the Monte  
2712 Carlo being produced assumed  $M_A^{QE} = 1.21\text{GeV}$  whilst the pre-fit dial value should  
2713 be  $M_A^{QE} = 1.03\text{GeV}$ , as suggested by [204]. Furthermore, the predicted event rates  
2714 of samples which target CCRES interaction modes is significantly reduced when  
2715 considering the post-ND fit. This follows the observations in subsection 6.4.3. The  
2716 strength of the accelerator neutrino experiment can also be seen in the remarkable  
2717 difference between the oscillated and unoscillated predictions in the FHC1Rmu and  
2718 RHC1Rmu samples. There is a very obvious decrease in the expected event rate between  
2719 the two predictions which is not as clearly observed in the atmospheric samples.  
2720 This is due to the fact that the beam energy is tuned to the maximum disappearance  
2721 probability, which is not the case for the naturally generated atmospheric neutrinos.



**Figure 8.2.:** Two-dimensional log-likelihood scan of the appearance ( $\sin^2(\theta_{13})$  -  $\delta_{CP}$ ) parameters showing the response of the beam samples (top), atmospheric samples (middle) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are assumed to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.



**Figure 8.3.:** Two-dimensional log-likelihood scan of the disappearance ( $\sin^2(\theta_{23})$  -  $\Delta m_{32}^2$ ) parameters showing the response of the beam samples (top), atmospheric samples (middle) and the summed response (bottom). The Asimov A oscillation parameters, defined in Table 2.2, are assumed to be the true point (Black Cross). The position of the smallest log-likelihood is highlighted with the triangle. Prior uncertainty terms of the oscillation parameters are neglected. The two(three) sigma contour levels are illustrated with the dashed(solid) red line.

Sample	Total Predicted Events					
	Generated		Pre-fit		Post-fit	
	Osc	UnOsc	Osc	UnOsc	Osc	UnOsc
SubGeV- <i>elike</i> -0d <sub>c</sub> y	7121.0	7102.6	6556.8	6540.0	7035.2	7015.7
SubGeV- <i>elike</i> -1d <sub>c</sub> y	704.8	725.5	693.8	712.8	565.7	586.0
SubGeV- <i>mulike</i> -0d <sub>c</sub> y	1176.5	1737.2	1078.6	1588.1	1182.7	1757.1
SubGeV- <i>mulike</i> -1d <sub>c</sub> y	5850.7	8978.1	5351.7	8205.1	5867.0	9009.9
SubGeV- <i>mulike</i> -2d <sub>c</sub> y	446.9	655.2	441.6	647.7	345.9	505.6
SubGeV- <i>pi0like</i>	1438.8	1445.4	1454.9	1461.1	1131.1	1136.2
MultiGeV- <i>elike</i> - <i>nue</i>	201.4	195.6	201.1	195.3	202.6	196.7
MultiGeV- <i>elike</i> - <i>nuebar</i>	1141.5	1118.3	1060.7	1039.5	1118.5	1095.7
MultiGeV- <i>mulike</i>	1036.7	1435.8	963.1	1334.1	1015.2	1405.9
MultiRing- <i>elike</i> - <i>nue</i>	1025.1	982.2	1026.8	984.3	1029.8	986.4
MultiRing- <i>elike</i> - <i>nuebar</i>	1014.8	984.5	991.0	962.0	1008.9	978.5
MultiRing- <i>mulike</i>	2510.0	3474.4	2475.6	3425.8	2514.6	3480.4
MultiRingOther-1	1204.5	1279.1	1205.8	1280.3	1207.4	1281.0
PCStop	349.2	459.2	338.4	444.7	346.8	456.1
PCThr	1692.8	2192.5	1661.5	2149.8	1689.2	2187.8
UpStop- <i>mu</i>	751.2	1295.0	739.7	1271.6	750.4	1293.0
UpThruNonShower- <i>mu</i>	2584.4	3031.6	2577.9	3019.4	2586.8	3034.0
UpThruShower- <i>mu</i>	473.0	488.6	473.2	488.7	473.8	489.4
FHC1R <sub>mu</sub>	328.0	1409.2	301.1	1274.7	345.1	1568.0
RHC1R <sub>mu</sub>	133.0	432.3	122.7	396.2	135.0	443.9
FHC1Re	84.6	19.2	77.4	18.2	93.7	19.7
RHC1Re	15.7	6.4	14.6	6.1	15.9	6.3
FHC1Re1de	10.5	3.2	10.3	3.1	8.8	2.9

**Table 8.1.:** The Monte Carlo prediction of each sample observed at SK used within this analysis. Three model parameter tunes are considered, as defined in subsection 6.4.3. The oscillated predictions assumed Asimov A oscillation parameters provided in Table 2.2.

2723 **Appendix A**

2724 **Atmospheric Sample Spectra**

2725 This appendix documents the interaction mode breakdown of all the atmospheric  
2726 samples used within the analysis. The generated tune of the model parameters and the  
2727 Asimov A oscillation parameter set (defined in Table 2.2) are assumed. The livetime of  
2728 SK-IV is taken to be 3244.4 days.

2729 **A.1. Binning**

2730 The lepton momentum and cosine zenith binning edges for the atmospheric samples  
2731 used within this analysis are defined in Table A.1.

2732 **A.2. Fully Contained Sub-GeV Samples**

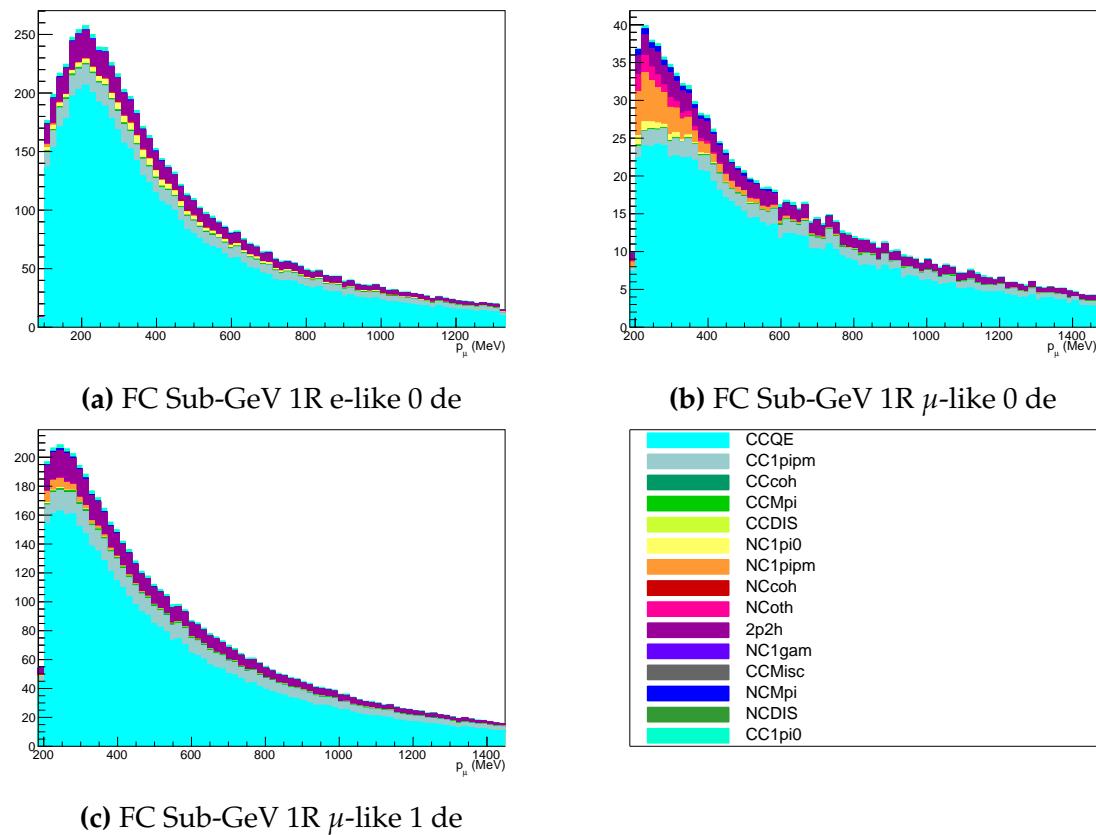
2733 The interaction mode breakdown of the fully contained Sub-GeV samples are shown  
2734 in Figure A.1 and Figure A.2, for the samples with enriched CC0 $\pi$  and CC1 $\pi^\pm$   
2735 respectively.

2736 The CC0 $\pi$  sample are dominated by CCQE events ( $\sim 70\%$ ) with smaller contributions of 2p2h ( $\sim 12\%$ ) and CC1 $\pi$  ( $\sim 10\%$ ) components. The energy peaks around 300 MeV, which is slightly below that of the T2K samples but still has significant contribution upto 1 GeV which overlaps the T2K sample energy range.

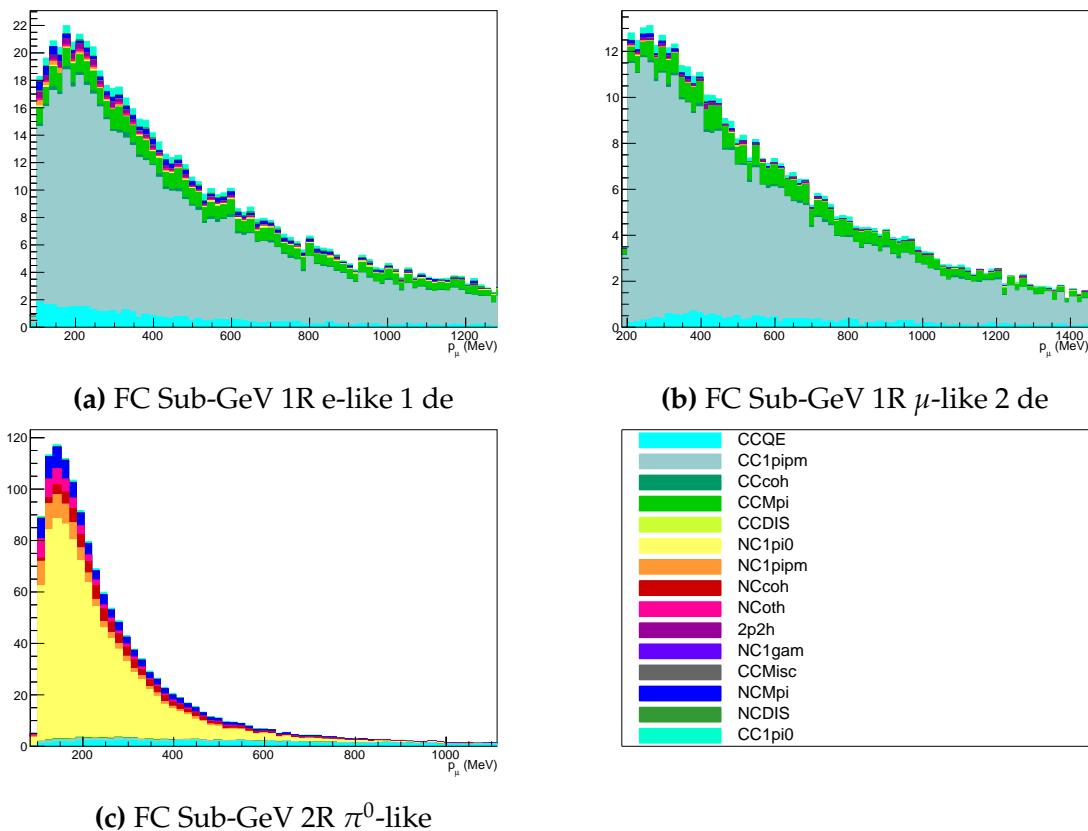
2740 The one-ring CC1 $\pi$  samples, where the pion is tagged via its decay electron, are  
2741 dominated by CC1 $\pi$  events ( $\sim 75\%$ ) with a small contribution of CCM $\pi$  ( $\sim 10\%$ ).  
2742 The two-ring pion sample is mostly dominated by the NC1 $\pi^0$  via resonances, and  
2743 has several equally-sized contributions from CCQE, NC1 $\pi^\pm$  via resonances, and NC  
2744 coherent pion production, where the  $\pi^0$  likely comes from nucleon and  $\pi^\pm$  final state  
2745 interactions in the nucleus.

Sample	$\cos(\theta_Z)$ Bins	Momentum Bin Edges ( $\log_{10}(P)$ MeV)
SubGeV- <i>elike</i> -0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV- <i>elike</i> -1dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV- <i>mulike</i> -0dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV- <i>mulike</i> -1dcy	10	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV- <i>mulike</i> -2dcy	1	2.0, 2.4, 2.6, 2.8, 3.0, 3.2
SubGeV- <i>pi0like</i>	1	2.0, 2.2, 2.4, 2.6, 2.8, 3.2
MultiGeV- <i>elike</i> -nue	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV- <i>elike</i> -nuebar	10	3.0, 3.4, 3.7, 4.0, 5.0
MultiGeV- <i>mulike</i>	10	3.0, 3.4, 5.0
MultiRing- <i>elike</i> -nue	10	3.0, 3.4, 3.7, 5.0
MultiRing- <i>elike</i> -nuebar	10	3.0, 3.4, 3.7, 5.0
MultiRing- <i>mulike</i>	10	2.0, 3.124, 3.4, 3.7, 5.0
MultiRing- <i>Other1</i>	10	3.0, 3.4, 3.7, 4.0, 5.0
PC-Stop	10	2.0, 3.4, 5.0
PC-Through	10	2.0, 3.124, 3.4, 3.7, 5.0
Upmu-Stop	10	3.2, 3.4, 3.7, 8.0
Upmu-Through-Showering	10	2.0, 8.0
Upmu-Through-NonShowering	10	2.0, 8.0

**Table A.1.:** The reconstructed cosine zenith and lepton momentum binning assigned to the atmospheric samples. The “ $\cos(\theta_Z)$  Bins” column illustrates the number of bins uniformly distributed over the  $-1.0 \leq \cos(\theta_Z) \leq 1.0$  region for fully and partially contained samples and  $-1.0 \leq \cos(\theta_Z) \leq 0.0$  region for up- $\mu$  samples.



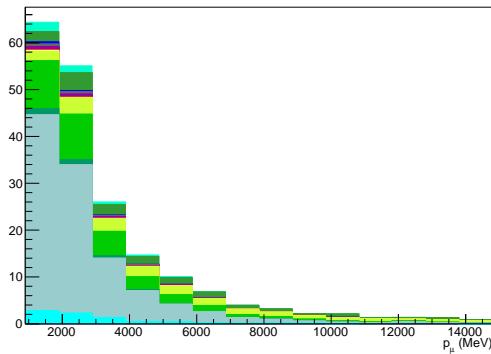
**Figure A.1:** Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting CC $0\pi$  events.



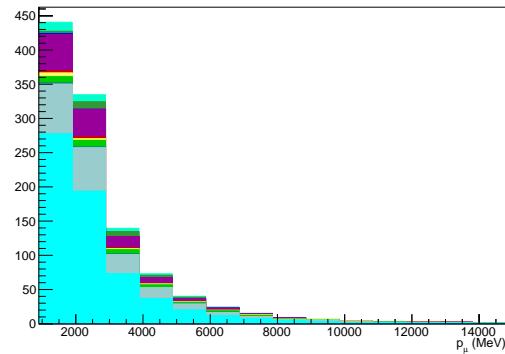
**Figure A.2:** Breakdown by interaction mode of the FC Sub-GeV atmospheric samples targeting single pion events.

### **A.3. Fully Contained Multi-GeV Samples**

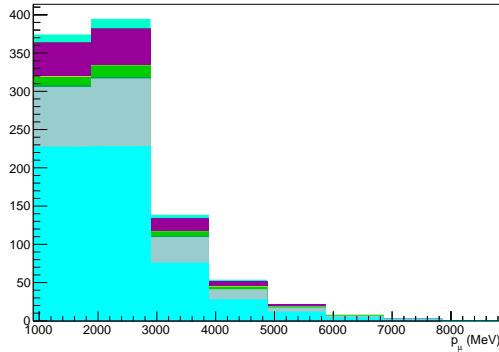
The interaction mode breakdown of fully contained multi-GeV samples is highlighted in Figure A.3. Due to the event selection applied in SK which targets  $\pi^+$  and  $\pi^-$  separation, the  $\nu_e$  sample mainly consists of events with pions (single pion production or multi-pion/DIS interactions). The pion separation is explained in Section section 6.1. This reasoning also explains the significant CCQE contribution of the  $\bar{\nu}_e$  sample. The muon-like sample is dominated by CCQE interactions with  $\sim 10 - 15\%$  2p2h and CC1 $\pi$  contribution of events.



(a) FC Multi-GeV single ring  $\nu_e$ -like



(b) FC Multi-GeV single ring  $\bar{\nu}_e$ -like

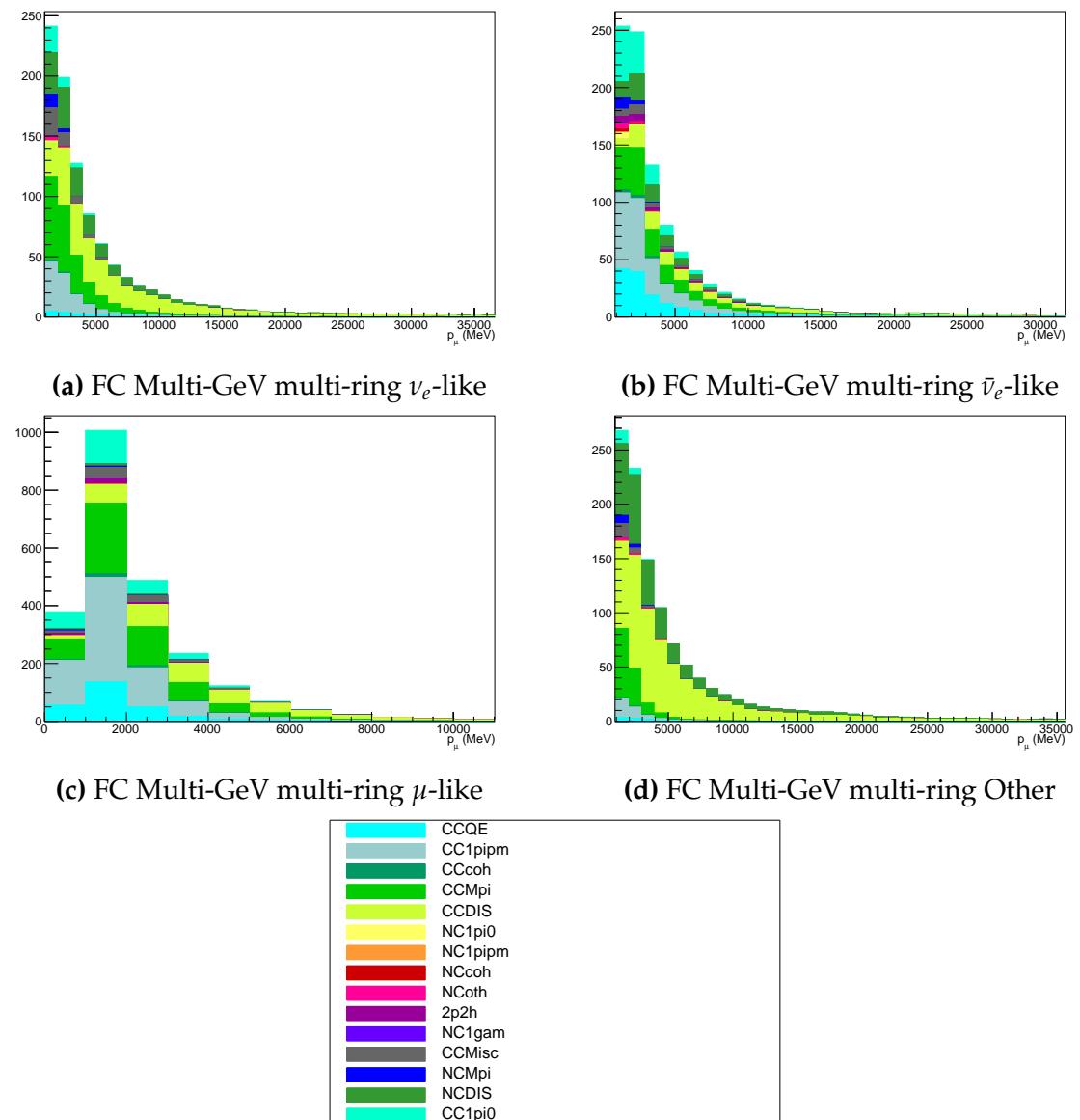


(c) FC Multi-GeV single ring  $\mu$ -like

**Figure A.3:** Breakdown by interaction mode of the FC Multi-GeV single ring atmospheric samples.

## <sup>2754</sup> A.4. Fully Contained Multi-Ring Samples

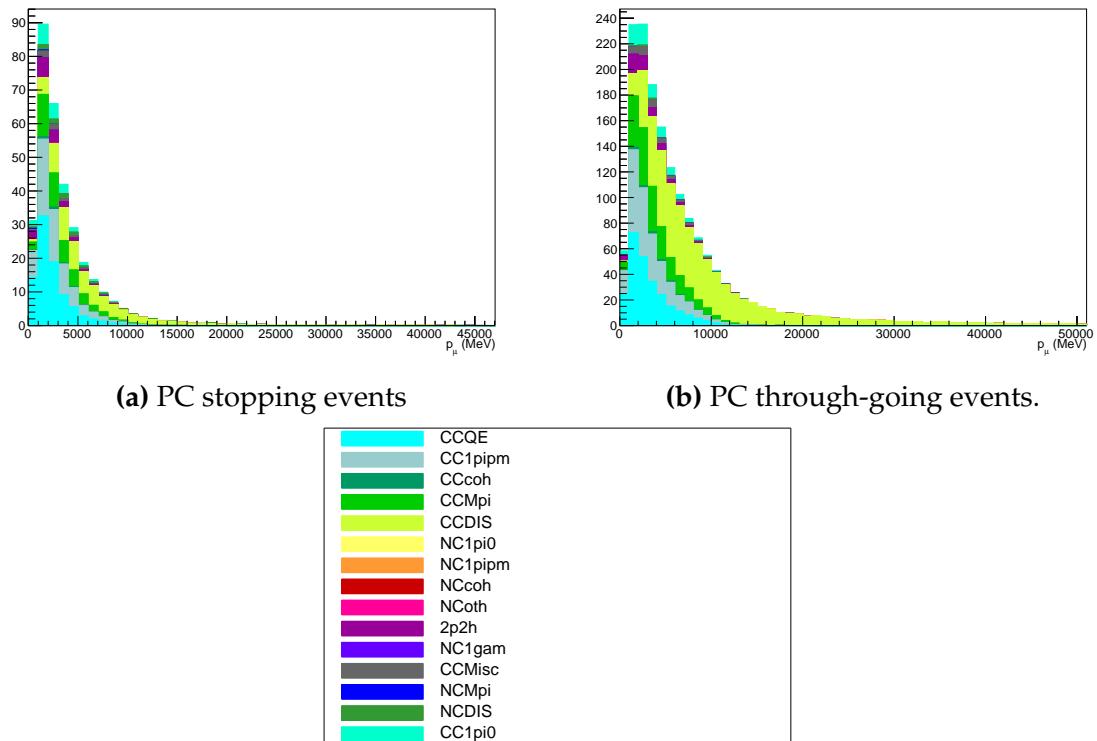
<sup>2755</sup> The interaction mode breakdown of fully contained multi-ring events is shown in Figure A.4. These samples see more interaction modes contributing in general, and there  
<sup>2756</sup> is a much larger contribution from multi-pion and DIS interaction modes, compared  
<sup>2757</sup> to the other samples.



**Figure A.4:** Breakdown by interaction mode of the FC Multi-GeV multi-ring atmospheric samples.

## <sup>2759</sup> A.5. Partially Contained Samples

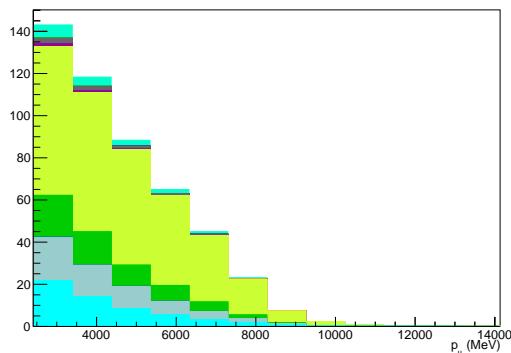
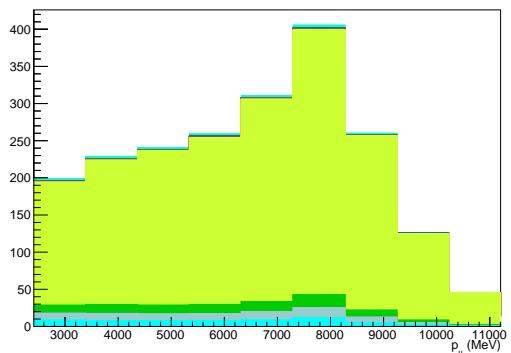
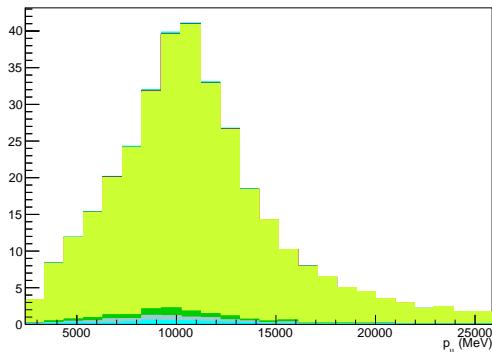
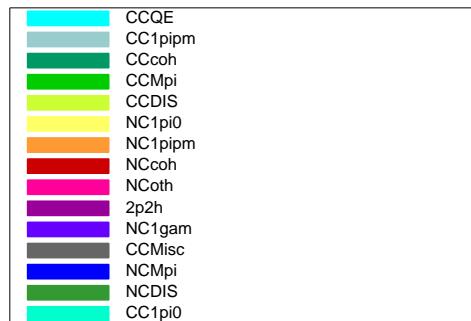
<sup>2760</sup> The breakdown for partially contained samples is highlighted in Figure A.5. As with  
<sup>2761</sup> the multi-ring samples, there is no dominating interaction mode. The neutrino energies  
<sup>2762</sup> of events in this sample extend into the tens of GeV and become dominated by DIS  
<sup>2763</sup> interaction modes in the high energy limit.



**Figure A.5.:** Breakdown by interaction mode of the PC atmospheric samples.

## <sup>2764</sup> A.6. Upward-Going Muon Samples

<sup>2765</sup> The breakdown for upward-going muons is illustrated in Figure A.6. These samples  
<sup>2766</sup> are significantly dominated by DIS interactions with energies extending up into the  
<sup>2767</sup> hundreds of GeV.

(a) Up- $\mu$  stopping events(b) Up- $\mu$  through going non showering events(c) Up- $\mu$  through going showering events**Figure A.6.:** Breakdown by interaction mode of the atmospheric upward going muon samples.

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