





- Simulated events and rate results, along with input configurations, are systematically archived for easy access and future analysis. Additionally, all interpolators used in the process are preserved for future applications.
- Most cosmology-related calculations within the *ler* package are performed using the *astropy* library (Astropy Collaboration 2013). The default cosmological model is LambdaCDM ( $H_0 = 70$ ,  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$ ); however, users have the flexibility to employ any cosmology available in *astropy*. All internal calculations in *ler* will then be based on the user-selected cosmological model.

## Equations

**Detectable Unlensed rates:**

$$R_U = \int dz_s \frac{dV_c}{dz_s} \frac{R_m(z_s)}{1+z_s} \{ \Theta[\rho(z_s, \theta) - \rho_{th}] P(\theta) d\theta \}$$

$z_s$ : GW source redshift,  $\frac{dV_c}{dz_s}$ : Differential co-moving volume,  $\frac{1}{1+z_s}$ : Time dilation correction factor,  $R_m(z_s)$ : source frame merger rate density,  $\theta$ : GW source parameters,  $P$ : probability distribution,  $\rho$ : SNR,  $\rho_{th}$ : SNR threshold,  $\Theta$ : Heaviside function to select detectable GW events.

**Detectable Lensed rates:**

$$R_L = \int dz_s \frac{dV_c}{dz_s} \tau(z_s) \frac{R_m(z_s)}{1+z_s} \mathcal{O}_{images}(z_s, \theta, \mu_i, \Delta t_i, \rho_{th}) P(\theta) P(\theta_L | SL, z_s) P(\beta | SL) d\theta d\beta d\theta_L dz_s$$

$\tau(z_s)$ : Optical-depth of strong lensing,  $\theta_L$ : lens parameters,  $\beta$ : source position,  $\mu$ : image magnification,  $\Delta t$ : image time delay,  $\mathcal{O}$ : operator to select detectable lensed GW events,  $i$ : index of images of a lensed event, SL: strong lensing condition.

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