

# New seismological constraints on the available parameter space of macroscopic dark matter

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Limits have been proposed constraining the available parameter space of macroscopic dark matter (Macros). These limits were based on the energy deposition of Macros into low frequency seismic modes being akin to that of a nuclear or chemical bomb. We show that this assumption does not hold by considering the mechanism through which Macro impacts generates seismic waves. Our new estimates take into account the effect of seismometer sensitivity and energy distribution across frequencies, anelastic attenuation, geometric effects, internal anisotropy of the Earth or Moon, as well as energy loss due to melting. We produce limits describing the Macro parameter space which can be ruled out by current seismological evidence.

## I. INTRODUCTION

Teplitz *et al.* [1] have suggested that the energy deposition into the 1 Hz range from a nuclearite (Macro) impact should be 0.05 of the total energy deposition. This, as shown in section II A, is a gross overestimate even compared to our own generous overestimation. We produce a more accurate model of the seismic effects of Macro impacts, including in our model the effects of geometric lensing, stratification of the Moon, anelastic attenuation, and geometric attenuation. The result of this analysis is an upper bound on the event rate that would be measured by the Apollo lunar seismometers. We also include a simple argument which demonstrates that a large portion of the available parameter space cannot be ruled out by seismological constraints.

The moon is an attractive site to carry out a seismological search for Macro impacts. Unlike the Earth, the Moon is nearly dormant. Most internal seismological activity in the Moon originates from deep Moonquakes, which are very attenuated by the time they reach the surface. Most of the noise in lunar seismograms comes from asteroids, but these are for the most part easily identifiable as such. Thus, in the case that we can distinguish between asteroids and Macros, the Moon is a nearly ideal location to place constraints via total seismic event rate. However, in the case that the Macro impacts are very weak, it may be impossible to distinguish them from asteroids or weak Moonquakes based on a single seismogram. In this case, it is necessary to identify correlated events among lunar seismometers which correspond to a propagating source of seismic waves. This sort of analysis will be left to future work.

## II. THE SEISMIC SOURCE

### A. Seismic Wave Generation

The energy from the initial impact generates a column of melted rock. The super-heated rock, held at constant volume, generates a pressure which sources the seismic waves that can be measured at detectors. Macro's with velocity insignificantly changed by their encounter with the Moon source a pressure of the form

$$p = p_0 f(r) \chi_{(r,z) \in [a,b] \times [-\frac{\ell}{2}, \frac{\ell}{2}]}, \quad (1)$$

where  $f$  is an analytic function whose range is a subset of  $[-1, 1]$ , and  $p_0$  is the maximum differential pressure in this region. After a short time, the pressure is still of this form, since the inhomogeneity of the Moon is significant only on large scales.

The seismometers on the moon have a net useful range of up to 20 Hz and are sensitive to displacements as small as 0.3 nm [2]. Therefore, we only calculate the energy contained in the modes of frequency less than 20 Hz.

The Fourier transform of  $p$  to leading order in  $kb$  is

$$P = \frac{4\pi p_0}{k \cos \theta} \sin\left(\frac{hk \cos \theta}{2}\right) F + \mathcal{O}((kb)^0), \quad (2)$$

where  $\theta, \phi$  are the momentum space polar angles,  $k$  is the magnitude of the wave-vector,  $F = b^2 f_1(b) - a^2 f_1(a)$  and

$$f_1(x) = \sum_{n=0}^{\infty} \sum_{m=0}^n \binom{n}{m} \frac{f^{(n)}(x_0)}{n!} \frac{(-x_0)^{n-m}}{2+m} x^m, \quad (3)$$

for some  $x_0 \in [a, b]$ . The energy of a seismic  $p$ -wave is

$$E = \int d^3x \left( \rho |\partial_t u|^2 + (\lambda + 2\mu) |\nabla u|^2 \right), \quad (4)$$

where  $u$  is the displacement field, and  $\rho, \lambda, \mu$  are the density and Lamé coefficients respectively. For  $p$ -waves,  $p = -K \nabla \cdot u$  where  $K$  is the bulk modulus. Define  $\kappa = K^2/(\lambda + 2\mu)$ . It follows that

$$E = \frac{1}{\kappa} \int \frac{d^3k}{(2\pi)^3} |P|^2. \quad (5)$$

Since a seismometer can only detect lower frequency modes, the detectible energy is

$$E_k = \frac{1}{\kappa} \int_0^k \frac{dk'}{(2\pi)^2} k'^2 \sin \theta |P|^2. \quad (6)$$

Substituting in our  $P$ , we find that

$$E_k = \frac{1}{\kappa} 2p_0^2 F^2 \ell k^2 \times \left( \frac{\sin(\ell k) + k\ell(\cos(\ell k) - 2)}{\ell^2 k^2} + \text{Si}(\ell k) \right) + \mathcal{O}((kb)^4).$$

When  $\ell k > 1$  it is a good approximation to take

$$E_k = \frac{1}{\kappa} \pi \ell p_0^2 F^2 k^2 := C_1 k^2. \quad (7)$$

Although the initial pressure is fixed by the fast behavior of the melt wave, linear elasticity may not be appropriate for the pressures near the source. As the pressure wave propagates outward, it is geometrically attenuated until the pressure differential is within the linear regime. The Macro travels at super-sonic speeds and will generate a shock wave. Shock waves typically evolve to pressure fronts resembling a right triangle [3] i.e.  $f = (r - r_0)/\Delta r$  where  $a = r_0$  and  $b = r_0 + \Delta r$ . This shape is also chosen for its simplicity. In this case  $f_1(x) = (2x - 3r_0)/6\Delta r$  and

$$E_k = \frac{\Delta r^2 (3r_0 + 2\Delta r)^2}{36\kappa} \pi \ell p_0^2 k^2. \quad (8)$$

The total energy of the pressure wave is

$$E = \frac{2\pi p_0^2 \ell}{\kappa} \int_a^b r dr f(r)^2, \\ = \frac{p_0^2}{6\kappa} \pi \Delta r \ell (4r_0 + 3\Delta r). \quad (9)$$

The Macro deposits energy

$$E_{\text{initial}} = \left| \frac{dE}{dx} \right| \ell = \rho \sigma_X v_X^2 \ell, \quad (10)$$

and loses some energy from phase changes of its surroundings. Thus, the pressure waves have initial energy

$$E_{\text{propagated}} = E_{\text{initial}} - E_{\text{melt}}, \\ = (\rho \sigma_X v_X^2 - \epsilon_{\text{melt}}) \ell := \epsilon \ell. \quad (11)$$

A very unrealistic assumption is that no energy is lost during the non-linear evolution. Nonetheless, it will produce a generous over-estimate of the energy in the linear regime. We thus set  $E = E_{\text{propagated}}$  and obtain an expression for  $r_0$

$$r_0 = \frac{3\kappa\epsilon}{2\pi p_0^2 \Delta r} - \frac{3\Delta r}{4}. \quad (12)$$

$r_0$  demarcates the end of the non-linear regime. The non-linear regime is characterized by faulting and fracturing. Brittle failure for granite occurs at stress exceeding  $3 \times 10^8$  Pa [4]. We thus take  $p_0 = \min(10^8 \text{ Pa}, p_{\text{source}})$ . Note that  $p_{\text{source}} \gg 10^8$  Pa for Macros that we will consider. By ignoring the effects of the non-linear regime on the seismic waves, we are grossly overestimating the energy that could reach detectors.

A lower bound on  $r_0$  is 0, corresponding to entirely linear behavior. This, in turn, imposes an upper bound on the pulse width

$$\Delta r \leq \sqrt{\frac{2\kappa\epsilon}{\pi p_0^2}}. \quad (13)$$

We now form  $\Xi := E_k/E_{\text{propagated}}$  representing the fraction of deposited energy detectable to seismometers

$$\Xi = k^2 \frac{(\pi p_0^2 \Delta r^2 - 18\kappa\epsilon)^2}{576\pi p_0^2 \kappa\epsilon}. \quad (14)$$

When  $\Delta r$  is restricted to its physical range,  $\Xi$  is a monotonic decreasing function in  $\Delta r$ . Thus

$$\frac{4}{9} \frac{\kappa\epsilon k^2}{\pi p_0^2} \leq \Xi < \frac{9}{16} \frac{\kappa\epsilon k^2}{\pi p_0^2}. \quad (15)$$

It is important to note that these expressions only hold for  $k(r_0 + \Delta r) \ll 1$ , however they will always provide an over-estimate of the fraction of detectable energy. Moreover, this condition holds for a wide variety of relevant parameters. Taking  $\epsilon_{\text{melt}} = 0$ ,  $\rho = 3.3 \times 10^3 \text{ kg m}^{-3}$ ,  $v_X = 2.5 \times 10^5 \text{ m s}^{-1}$ ,  $\kappa = 5.5 \times 10^{10} \text{ Pa}$ , and  $k = 1.5 \times 10^{-2} \text{ m}^{-1}$  [5], we obtain  $\Xi < 4.5 \times 10^4 (\sigma_X/\text{m}^2)$ . To obtain the figure that Teplitz *et al.* have used, one must take  $\sigma_X > 10^{-2} \text{ cm}^2$ , which is rather large.

It is important to note that  $p_0 \rightarrow 0$  does not imply  $\Xi \rightarrow \infty$ , since  $p_0$  is implicitly a function of  $\epsilon$  when  $p_0 < 10^8$  Pa. In this case, the ratio  $\epsilon/p_0^2$  is always finite since  $\epsilon \propto p_0^2$ .

### III. SEISMIC WAVE PROPAGATION

The velocity of  $p$ -waves as a function of distance from the center  $r$  of the Earth and Moon (and presumably other spherical rocky celestial bodies) is of the form  $v = a^2 - b^2 r^2$  [5][6]. Using Snell's law, we obtain the differential equations for the trajectory of  $p$ -wave rays within such bodies

$$\begin{aligned} v^2 &= \dot{r}^2 + \frac{p_{\text{ray}}^2 v^4}{r^2}, \\ 0 &= \dot{\theta} \pm \frac{p_{\text{ray}} v^2}{r^2}. \end{aligned} \quad (16)$$

where  $r$  is the radial coordinate of the ray,  $\theta$  is the polar angle of the ray measured from the center of the body, and  $p_{\text{ray}}$  is the ray parameter, which is fixed along a ray trajectory. These equations can be integrated. In the case that  $v = a^2 - b^2 r^2$ , we obtain for some constants  $q$  and  $\theta_0$ .

$$\begin{aligned} r(t) &= \frac{a\sqrt{(q^2 e^{2abt} - b)^2 + 4a^2 b^4 p_{\text{ray}}}}{b\sqrt{(q e^{2abt} + b)^2 + 4a^2 b^4 p_{\text{ray}}}}, \\ \tan(\theta(t) - \theta_0) &= \frac{q e^{4abt} - b^2}{4ab^3 p_{\text{ray}}} + ab p_{\text{ray}}. \end{aligned} \quad (17)$$

*A priori* these are geodesics on the Poincaré disc of radius  $\frac{a}{b}$ , which is easily seen from the form of  $v$ .

The Earth and Moon are stratified. At each boundary, a ray will be reflected and transmitted. We assume that the reflection and transmission coefficients are frequency independent for simplicity.

The last effect to account for is anelastic attenuation, which does depend on frequency. Anelastic attenuation is characterized by the quality factor  $Q$ , which is, in general, dependent on  $r$ . For a given mode, the ratio of final to initial amplitude is

$$\exp \left[ -k \int_{t_0}^t dt' \frac{v(r(t'))}{2Q(r(t'))} \right] := \exp[-k \text{Att}]. \quad (18)$$

The VPREMOON Model [5] provides piecewise constant data for  $Q$ , so it is reasonable to subdivide the moon further into strata of different  $Q$ . For propagation within a given layer  $i$ , the factor Att is given by  $\Delta t_i v_i / Q_i$ , where we take  $v_i$  to be the average velocity within a given strata. For the case of the Earth and Moon, this is a good approximation since

$v$  doesn't change significantly within a given layer of constant  $Q$ . Thus, the amplitude of a ray can be computed by knowing the two numbers

$$\begin{aligned} \text{Ref} &= \prod_i \text{Ref}_i, \\ \text{Att} &= \sum_i \frac{\Delta t_i v_i}{Q_i}, \end{aligned} \quad (19)$$

where  $\text{Ref}_i$  are the reflection or transmission coefficients of boundary  $i$  on which the ray is incident.

### IV. SEISMIC WAVE DETECTION

Consider a  $p$ -wave traveling towards positive  $x$  and of compact support  $S$  in the plane normal to its motion. Denote  $A = \int_S dy dz$ . The displacement field is

$$u(\vec{x}, t) = \chi_{\mathbb{R} \times S} \int \frac{dk}{2\pi} U(k) e^{-ik(x - v_p t)}, \quad (20)$$

$$|u(\vec{x}, t)| \leq \int \frac{dk}{2\pi} |U(k)|, \quad (21)$$

and the energy is

$$E = \rho v_p^2 A \int \frac{dk}{2\pi} k^2 |U(k)|^2. \quad (22)$$

As before, we denote the energy in the low frequency spectrum

$$E_k = \rho v_p^2 A \int_0^k \frac{dk'}{2\pi} k'^2 |U(k')|^2. \quad (23)$$

It follows that

$$|U(k)| = \sqrt{\frac{2\pi}{\rho v_p^2 A}} \frac{1}{k} \sqrt{\frac{dE_k}{dk}}, \quad (24)$$

and from the estimate above

$$|u(\vec{x}, t)| \leq (2\pi \rho v_p^2 A)^{-1/2} \int_0^k \frac{dk'}{k'} \sqrt{\frac{dE_{k'}}{dk'}}. \quad (25)$$

In our case  $\Xi E = C_1 k^2$  before attenuation. After anelastic attenuation,  $E_k = \text{Ref} 2C_1 k e^{-k \text{Att}}$

$$|u(\vec{x}, t)| \leq \sqrt{\frac{2}{\rho v_p^2 A}} \sqrt{\frac{\text{Ref} \Xi E}{\text{Att}}} \frac{1}{k} \text{Erf} \left[ \sqrt{\frac{\text{Att} k}{2}} \right]. \quad (26)$$

(25) is impossible to integrate when  $E_k$  consists of the energies of two different rays, each with different

attenuation factors. If attenuation were important, we may average Att among the coincident rays and obtain an approximate expression. However, since we only wish to consider the lowest frequency modes,  $f < 20\text{Hz}$ , the suppression Att $k$  is order unity, and does not contribute significantly to  $u$ . Taking Att = 0, the displacement caused by the incident  $p$ -wave is bounded above by

$$|u(\vec{x}, t)| \leq \sqrt{\frac{4\text{Ref}\Xi E}{\pi k \rho v_p^2 A}}. \quad (27)$$

## V. SIMULATION

In previous works, e.g. [1], the homogeneous Earth and Moon models have been used to place constraints of the parameter space of macroscopic dark matter. This approach neglects lensing which occurs because of the velocity gradient and the spherical boundaries of strata, and neglects the energy losses from reflection and refraction across these boundaries. We use the analysis in section III to propagate rays, each carrying a fraction of the total energy deposited by the Macro, to the boundary of the Moon. We then use HEALPix to create an intensity map of the lunar surface for a representative sample of Macro impacts. We then convert the intensity maps to displacement maps, and using the sensitivity of the lunar seismometers, we obtain an average event rate that the lunar seismometers would measure.

### A. Data Generation

We consider a Macro trajectory which passes a distance  $D$  from the center of the moon. The Macro trajectory has length  $L$ , and along its trajectory we consider a set of  $M$  points. From each of these points we propagate  $N$  randomly oriented rays, each endowed with  $E_i = \rho(r_i)\sigma_X v_X^2 L/(MN)$  energy, where  $i \in \{1, \dots, M\}$  labels the point, and  $\rho$  is taken to be a function of the radius. Because of the spherical symmetry, each ray is confined to a subspace of two dimensions. We use the trajectories in section III to propagate the rays. Since the moon is stratified, we must halt the propagation at the boundary of each layer to reflect and transmit rays. During a ray's propagation through a given stratum, its time spent in that layer  $\Delta t_i$  is recorded, the attenuation factor in that layer Att $_i$  is recorded, and the reflection/refraction coefficient Ref $_i$  is recorded. Each time a ray makes contact with the surface of the

Moon, its position, Att, Ref, and  $\Delta t$  are recorded. The data produced by the propagation routine is a table of these values and the corresponding surface coordinates, as well as density at the source,  $\kappa$  at the source, and  $p$ -wave velocity at the source. In this model, it takes 16 iterations to propagate a ray from one side of the moon to the other. Rays which propagate from one side of the moon to the other only after 16 iterations will be weak from undergoing 16 refractions. Thus in order to save on computation time, it is reasonable to reduce the number of iterations to 9.

### B. Data Analysis

Analysis begins by converting the position data to HEALPix coordinates for a given  $n_{\text{side}}$ . We pick  $n_{\text{side}} = 4$  since this corresponds to an angular resolution of 7 degrees, which accounts for the effects of diffraction about the strata boundaries. Define  $A_{\text{pix}} = 4\pi R_{\text{moon}}^2/n_{\text{pix}}$  the area of a HEALPix pixel, where  $n_{\text{pix}} = 12 \times 2^{n_{\text{side}}}$ . The time it takes a ray to cross a HEALPix pixel can be estimated  $t_{\text{crossing}}^2 = A_{\text{pix}}/\pi v_{\text{surface}}^2$ . According to VPREF-MOON,  $v_{\text{surface}} = 1\text{ km s}^{-1}$ . Rays that contact a pixel within  $t_{\text{crossing}}$  of one another are considered to add constructively. Since the integral in section IV is unreasonable to do for more than one ray, we average Att over the coincident rays. We require that each ray distributes its energy evenly over the area of the HEALPix pixel, and thus set  $A = A_{\text{pix}}$ . This is a good approximation when  $N = n_{\text{pix}}$ , since the average angular separation of the rays at the source corresponds to the average angular separation of the HEALPix pixels.

Using (27) we compute an upper bound for the displacement within a given HEALPix pixel, and compare that to the sensitivity of the lunar seismometers given in [2]. If the calculated displacement is less than the seismometer sensitivity then the pixel is marked undetectable and is otherwise marked detectable. The number of detectable pixels is divided by the total number of pixels to obtain the fraction  $\mathcal{F}_D$  of the lunar surface with a detectable signal for a Macro making closest approach  $D$  to the lunar center.

This procedure is repeated for 17 values of  $D$  spaced evenly by 100 km. The average fraction of detectable surface area is then approximately

$$\mathcal{F} = \sum_D \frac{\mathcal{F}_D D}{D}. \quad (28)$$

If there are  $n$  lunar seismometers, we suppose that each occupies its own HEALPix pixel. Define  $m = \mathcal{F}n_{\text{pix}}$ , the average number of pixels with detectable frequency. If  $n + m \geq n_{\text{pix}}$ , then detection is guaranteed. Otherwise, if  $\mathcal{M}$  Macros impact the moon, the likelihood of any of them being detected is

$$\mathcal{P} = 1 - \left[ \frac{(n_{\text{pix}} - n)!(n_{\text{pix}} - m)!}{n_{\text{pix}}!(n_{\text{pix}} - n - m)!} \right]^{\mathcal{M}}. \quad (29)$$

### C. Homogeneous Limit

To obtain an estimate of the size of a detectable Macro, we suppose that one ray from each point source on the Macro Trajectory hits a single pixel. In the case that rays don't bend, this is a good approximation for behavior away from the source. Note, however, that this does not account for geometric focusing around the entrance and exit of the Macro trajectory.

$$u \leq \sqrt{\frac{\Xi E}{\pi^2 \rho k v_p^2 R_{\text{Moon}}^2}}. \quad (30)$$

This expression is independent of the choice of  $n_{\text{side}}$ . For the triangle wave,  $\sigma_X \leq 4.4 \times 10^{-7} \text{ cm}^2$  is unlikely to be detected unless it enters or exits the Moon adjacent to a seismometer.

## VI. RESULTS

Data has been generated but not run through the analysis routine yet. Preliminary results suggest that the computer analysis will not produce results much stronger than (30).

## VII. CONCLUSION

For Macros with density sufficient to pass nearly unimpeded through the Moon and sufficiently small

cross-section, and in particular nuclear dense objects such as strange quark nuggets with cross-section  $\sigma_X \leq 4 \times 10^{-7} \text{ cm}^2$ , we have demonstrated that seismological constraints are insufficient to rule them out. Since the Moon (and Earth) preferentially attenuate higher frequency modes, it is unlikely that measuring high frequencies would greatly improve sensitivity to these types of impacts. On the other hand, increasing the lunar seismometer sensitivity to small displacements may allow one to use seismological evidence to rule out high density Macros, although it is uncertain whether seismic noise at such small displacements would make Macro impacts indistinguishable from the background.

Future work will include new lower bounds on lunar and terrestrial seismic response to high density Macros. In the future, a more sensitive lunar array will allow for a better understanding of the lunar structure and for a greater volume of Macro parameter space to be considered through seismic means. In the case that the lunar seismic background is too loud for extremely small displacements to be distinguished, it will be necessary to reconstruct the seismic response to high density Macros in order to measure an approximate event rate.

## VIII. ACKNOWLEDGEMENTS

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## IX. SUMMER REFLECTION

My research this summer will be very important when I make my case to graduate schools that I am capable of doing independent research and problem solving. This project will likely be developed into my capstone, and we will try to publish this work. For next year's students, drink lots of coffee.

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- [1] W. B. Banerdt, T. Chui, E. T. Herrin, D. Rosenbaum, and V. L. Teplitz, in *The Identification of Dark Matter* (2005) pp. 581–586.
  - [2] G. Latham, M. Ewing, J. Dorman, Y. Nakamura, F. Press, N. Toksöz, G. Sutton, F. Duennebier, and D. Lammlein, *The Moon* **7**, 396 (1973).
  - [3] J. W. Forbes, *Shock wave compression of condensed matter: a primer* (Springer Science & Business Media, 2013).
  - [4] D. A. Lockner and N. M. Beeler, *International Geophysics* **81**, 505 (2002).
  - [5] R. F. Garcia, J. Gagnepain-Beyneix, S. Chevrot, and P. Lognonné, *Physics of the Earth and Planetary In-*

teriors **188**, 96 (2011).

[6] A. M. Dziewonski and D. L. Anderson, Physics of the earth and planetary interiors **25**, 297 (1981).