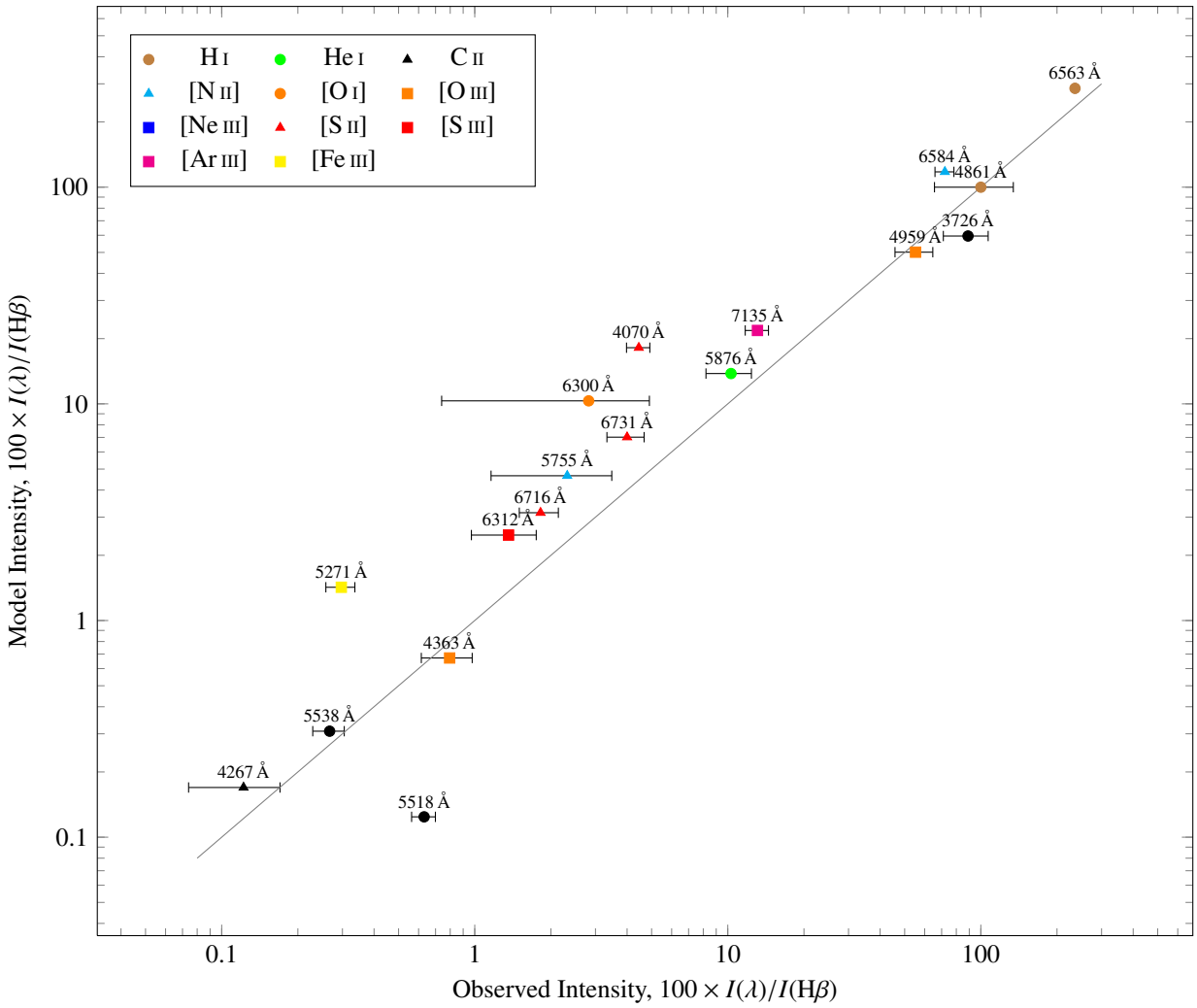
Model A: Baseline model

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set Cloudy Orion

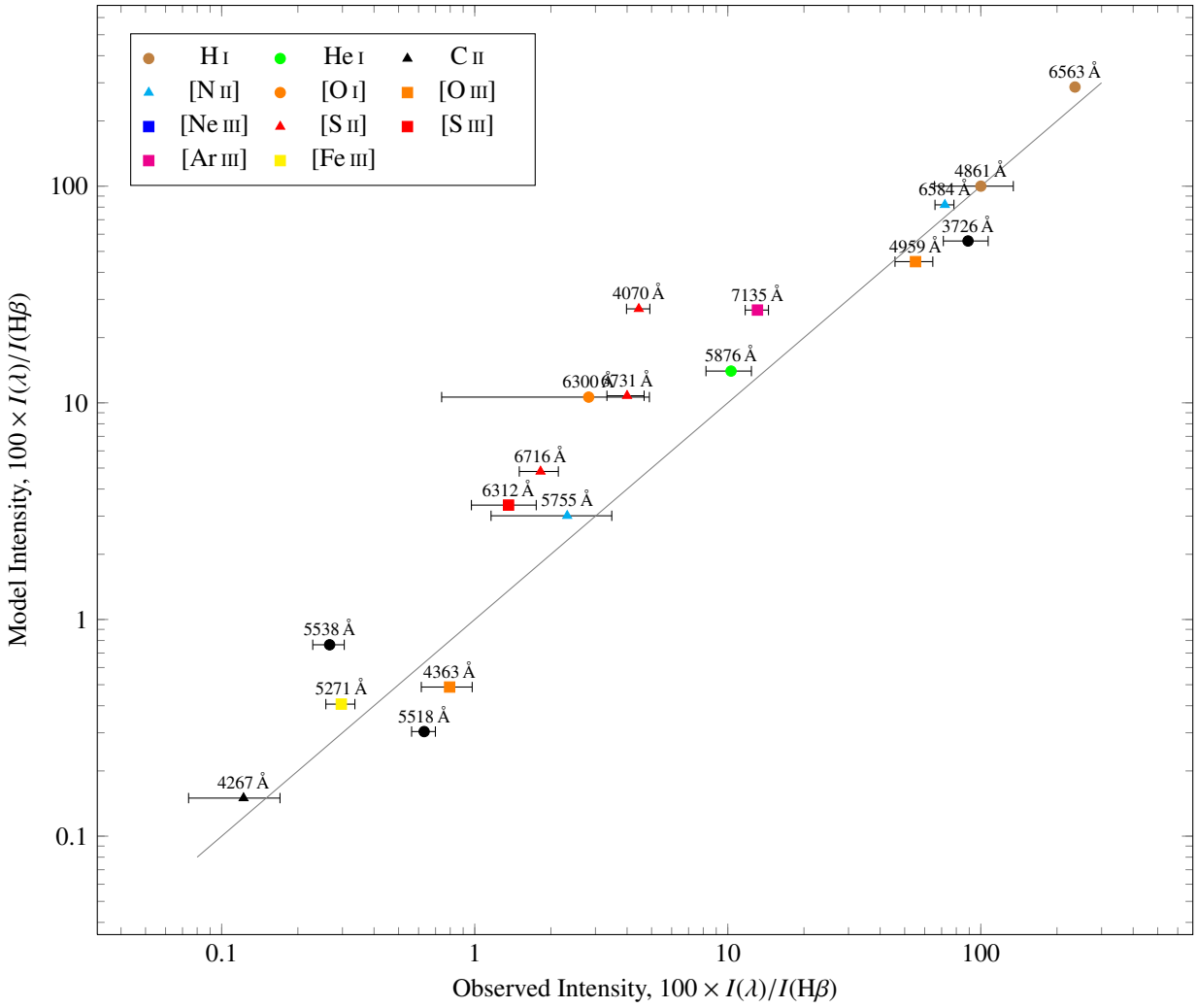


Model B: Esteban Set

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set Esteban et al. (2004), M42, $t^2 = 0.002$

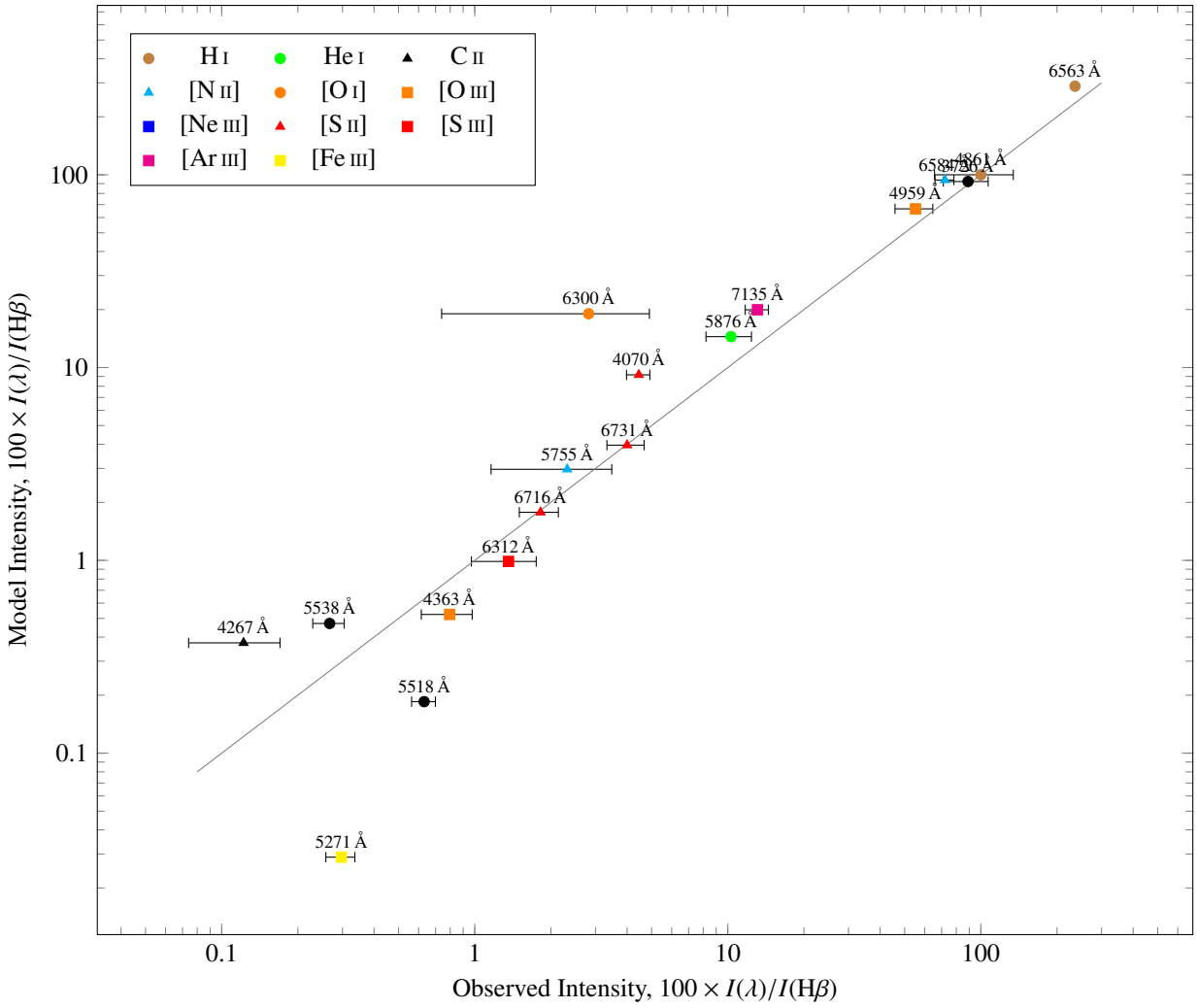


Model C: Tsamis for LV2

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set Tsamis et al. (2011), LV2

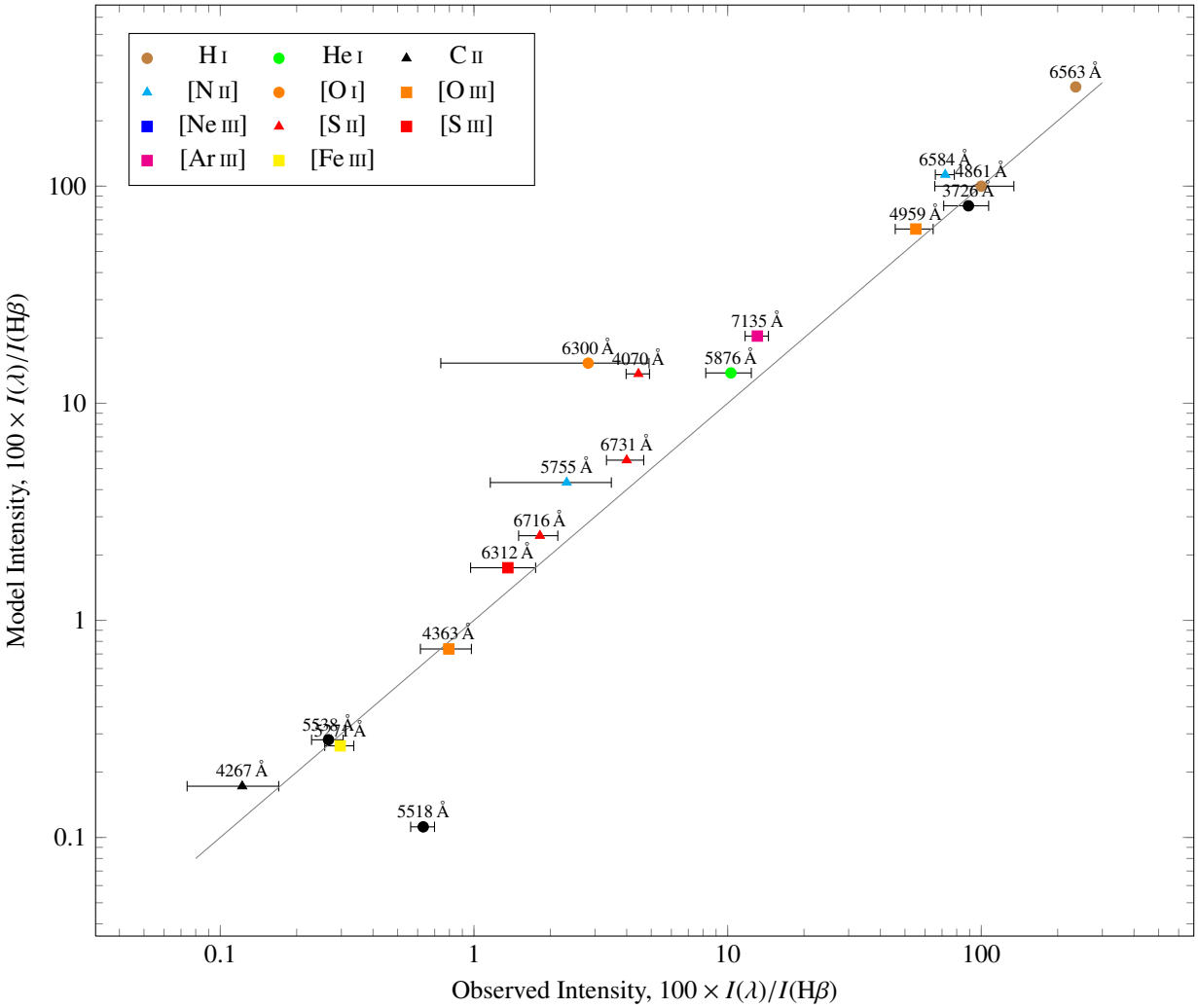


Model D: Tsamis HST10 preliminary

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set

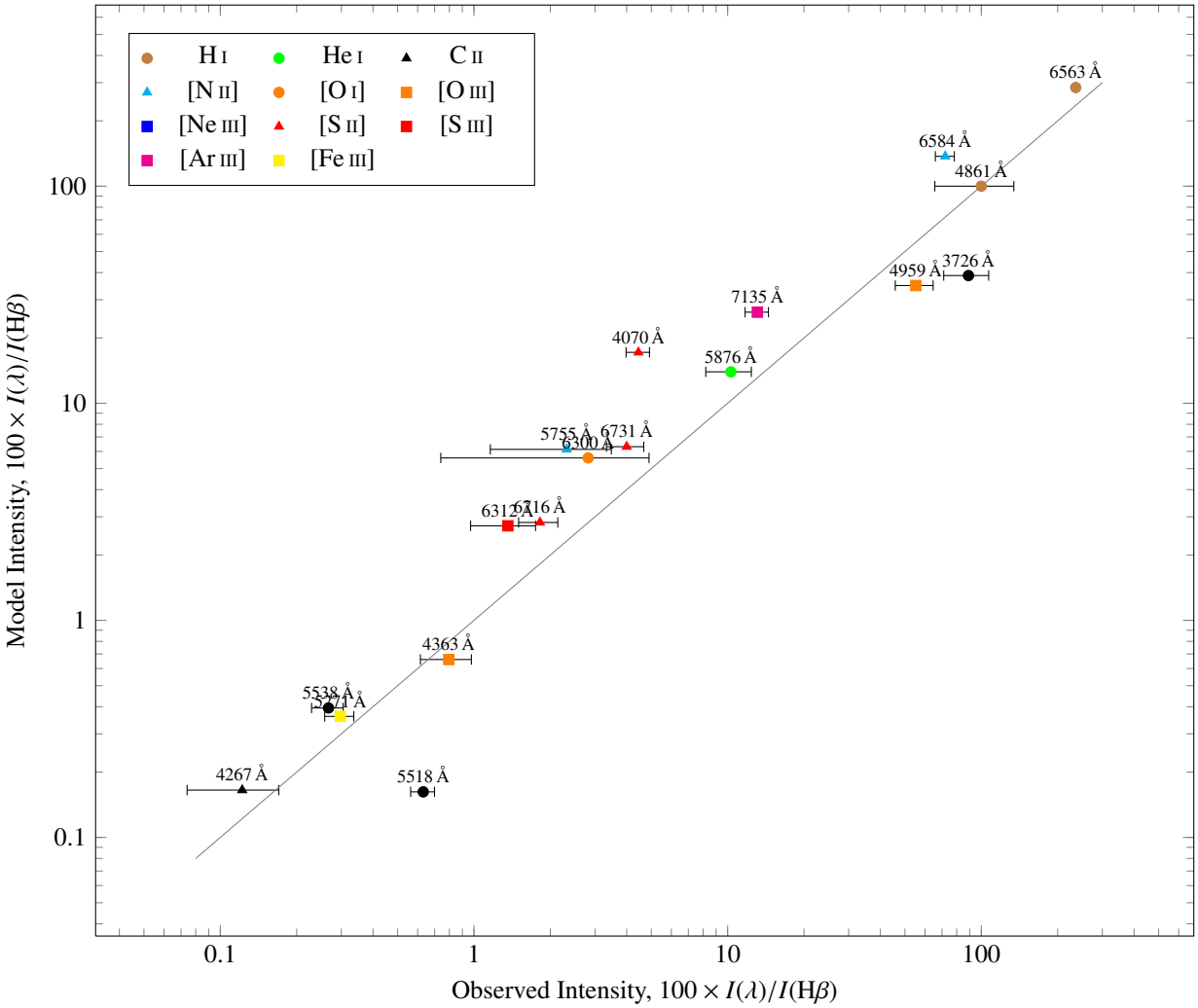


Model E: Tsamis HST10 new1

Spectrum Tlusty, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set

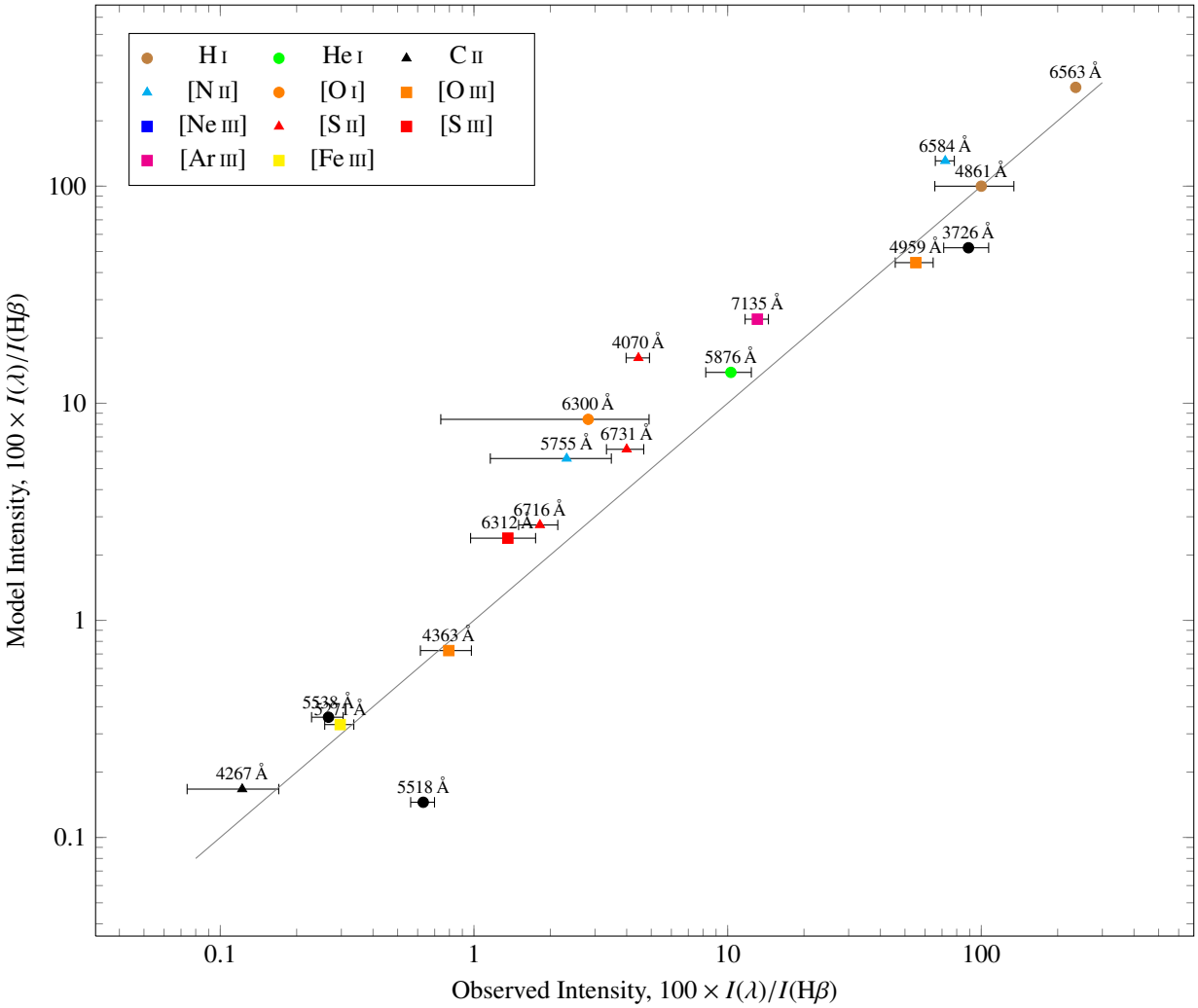


Model F: Tsamis HST10 new2

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set

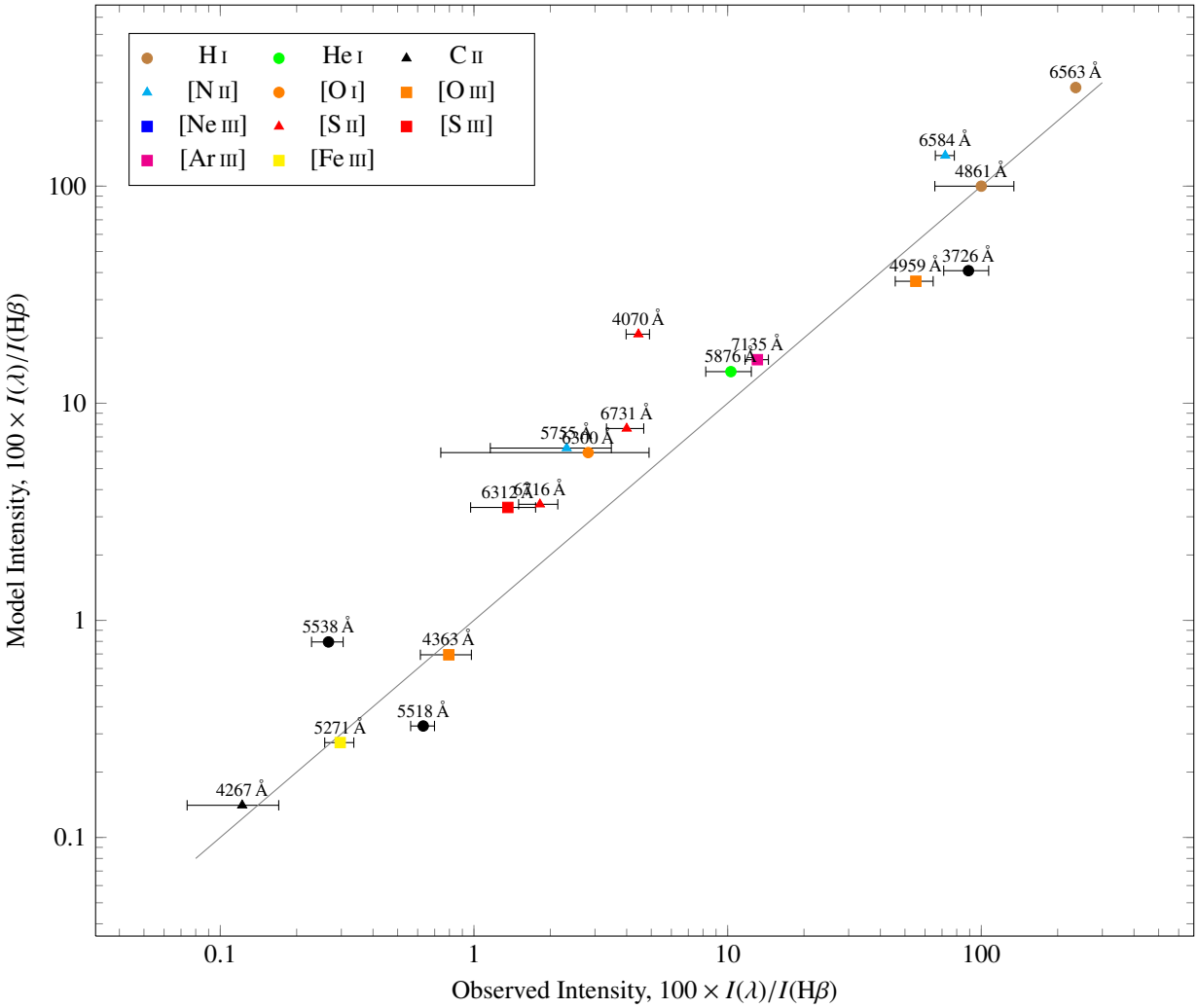


Model G: ZZ03

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set

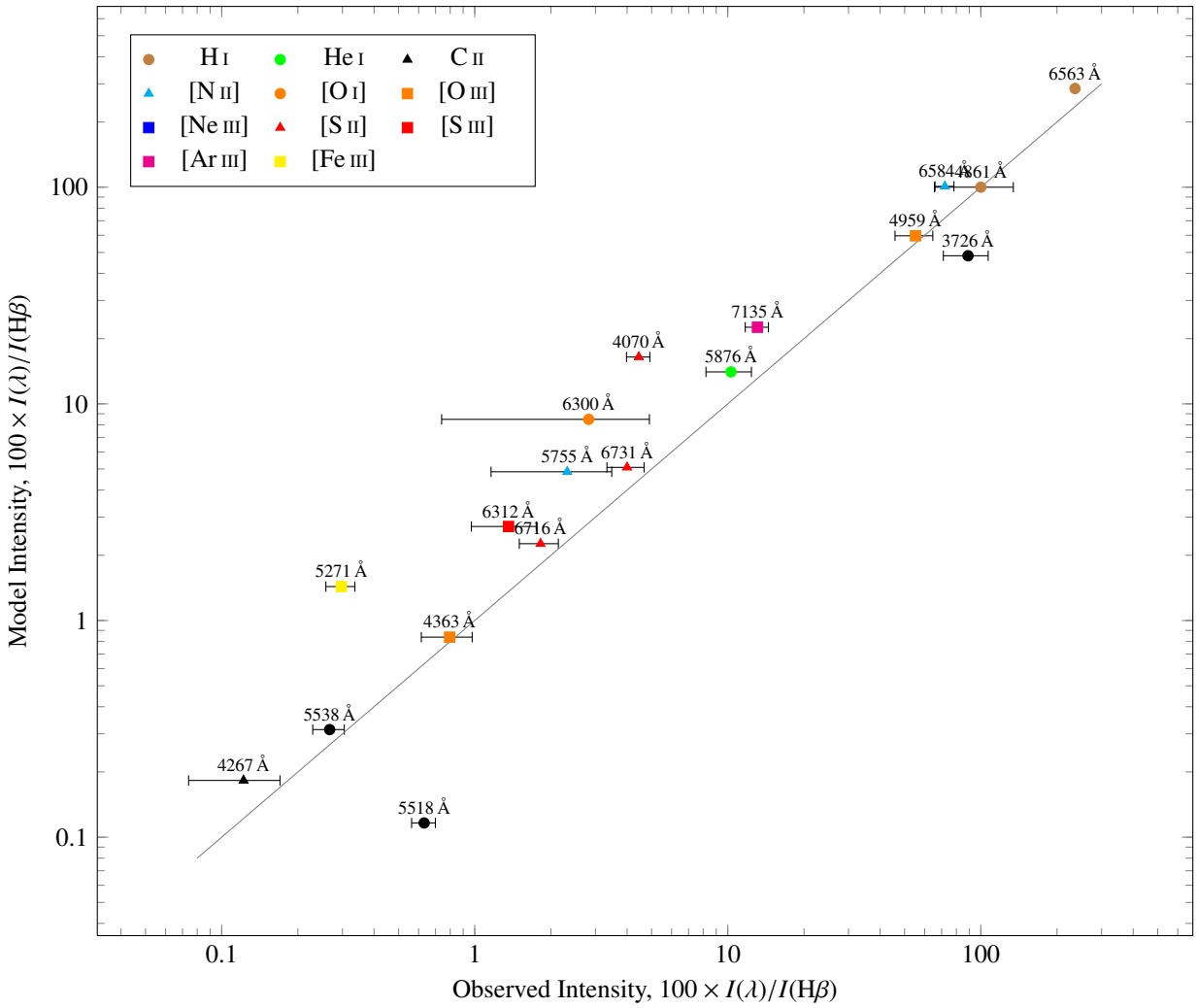


Model H: Baseline model

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set Cloudy Orion

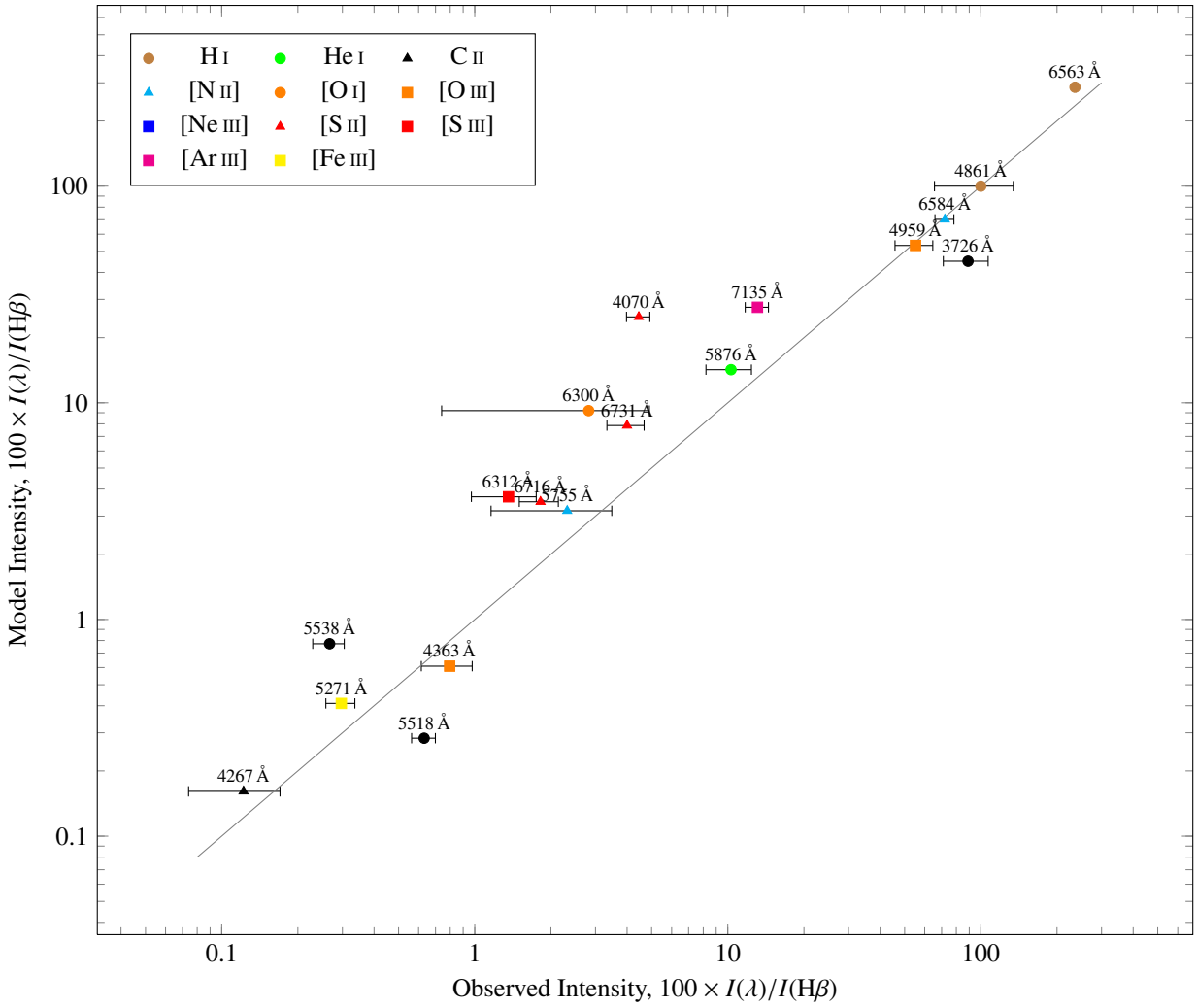


Model I: Esteban Set

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set Esteban et al. (2004), M42, $t^2 = 0.002$

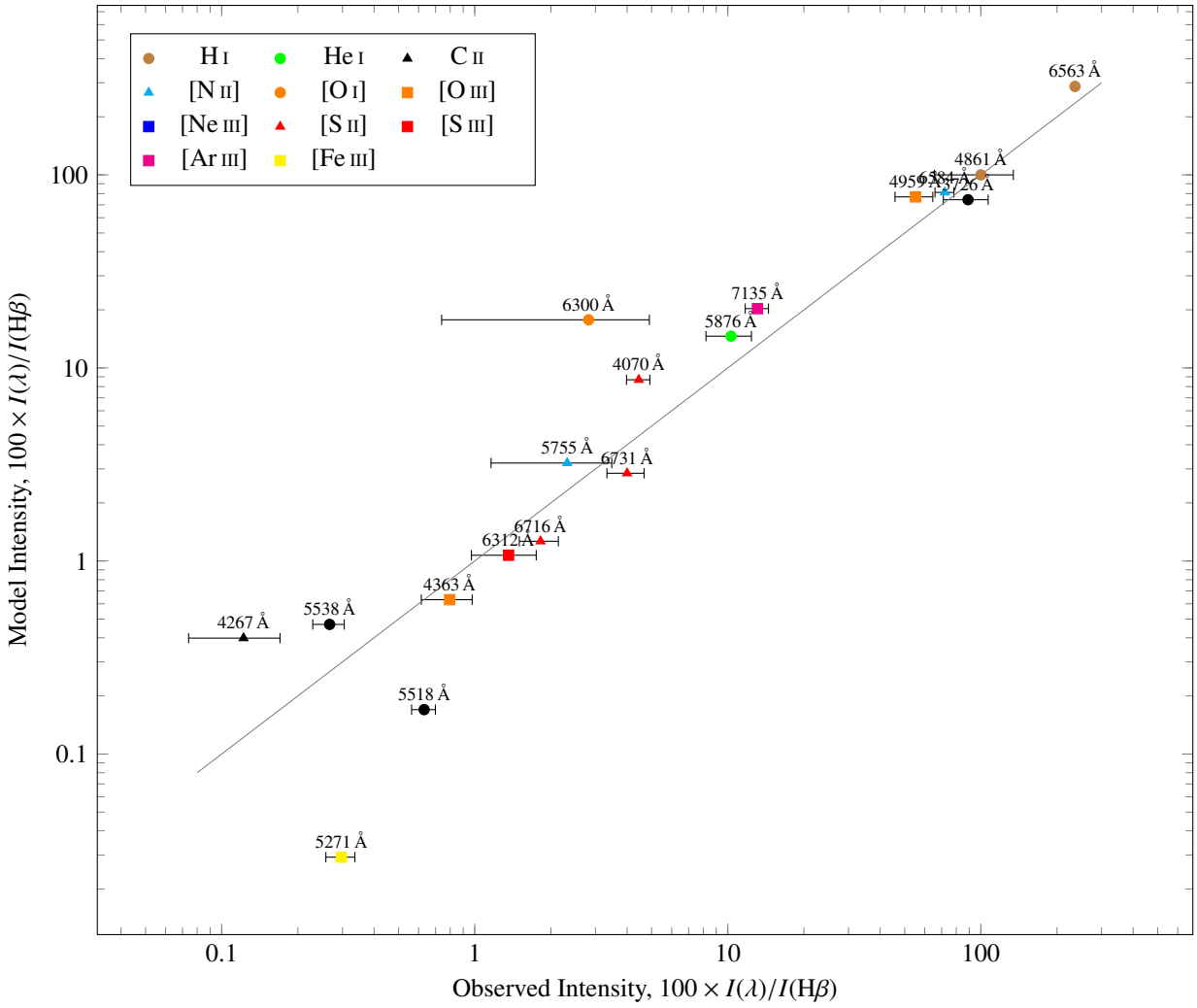


Model J: Tsamis for LV2

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set Tsamis et al. (2011), LV2

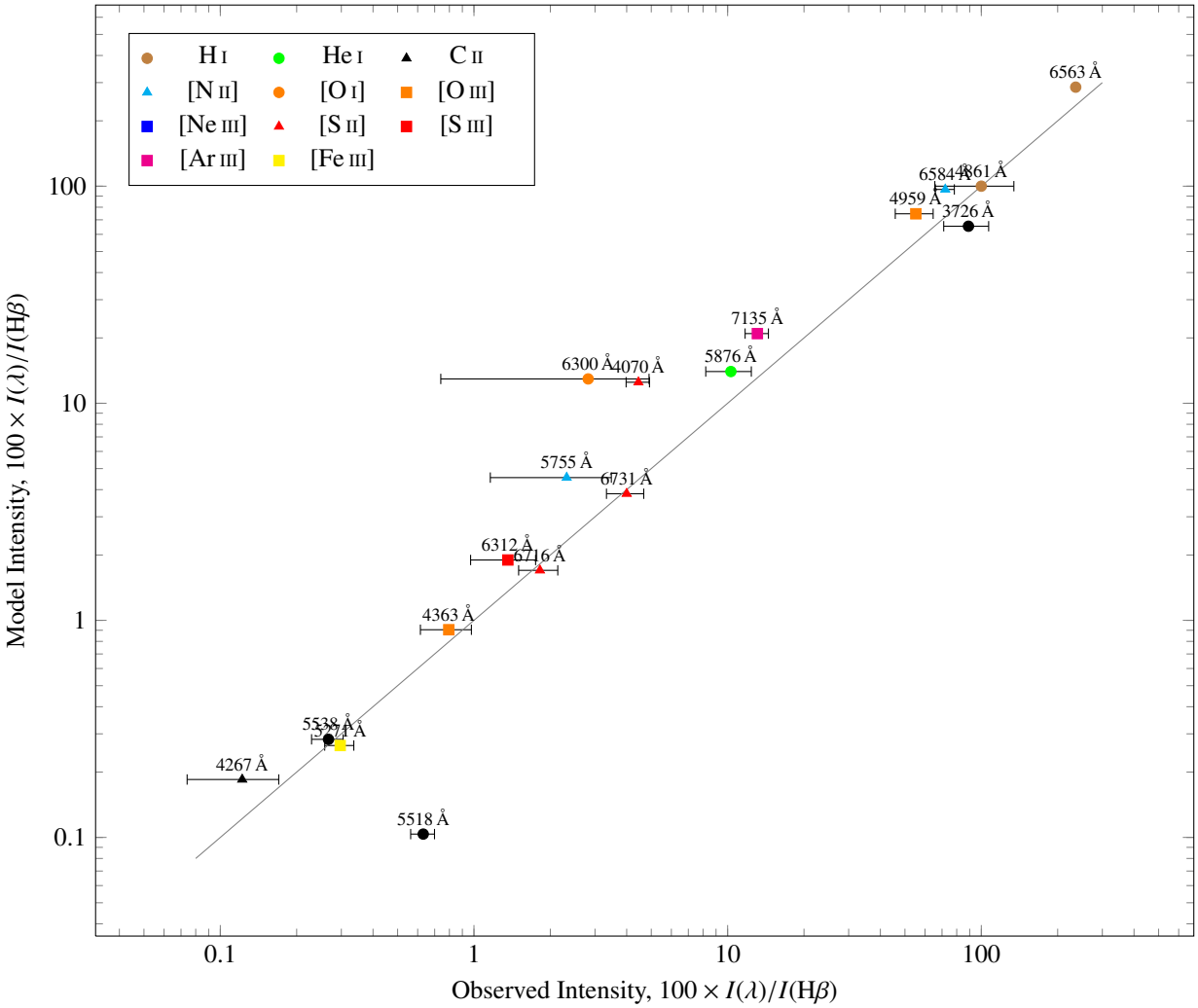


Model K: Tsamis HST10 preliminary

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set

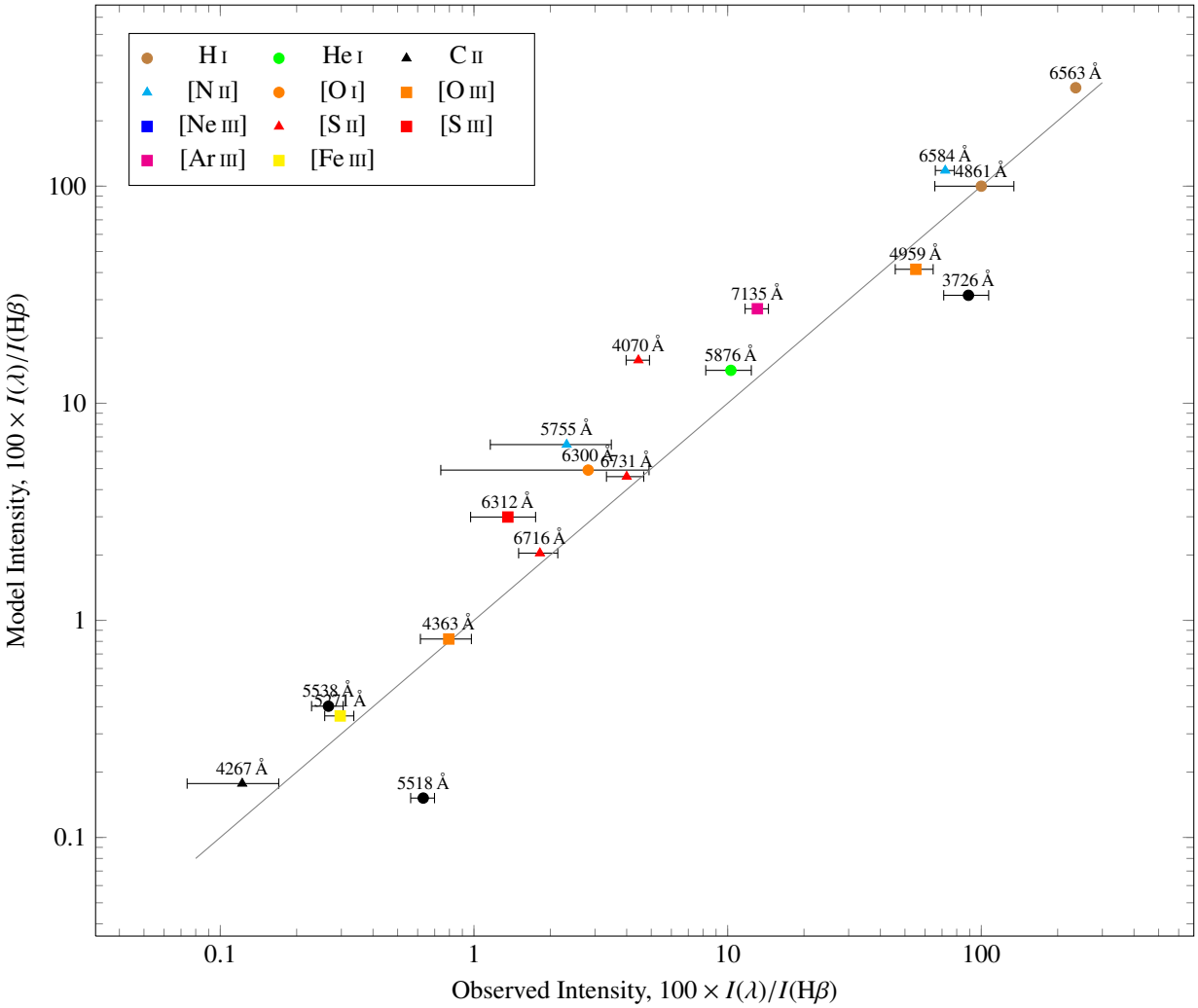


Model L: Tsamis HST10 new1

Spectrum Tlusty, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set

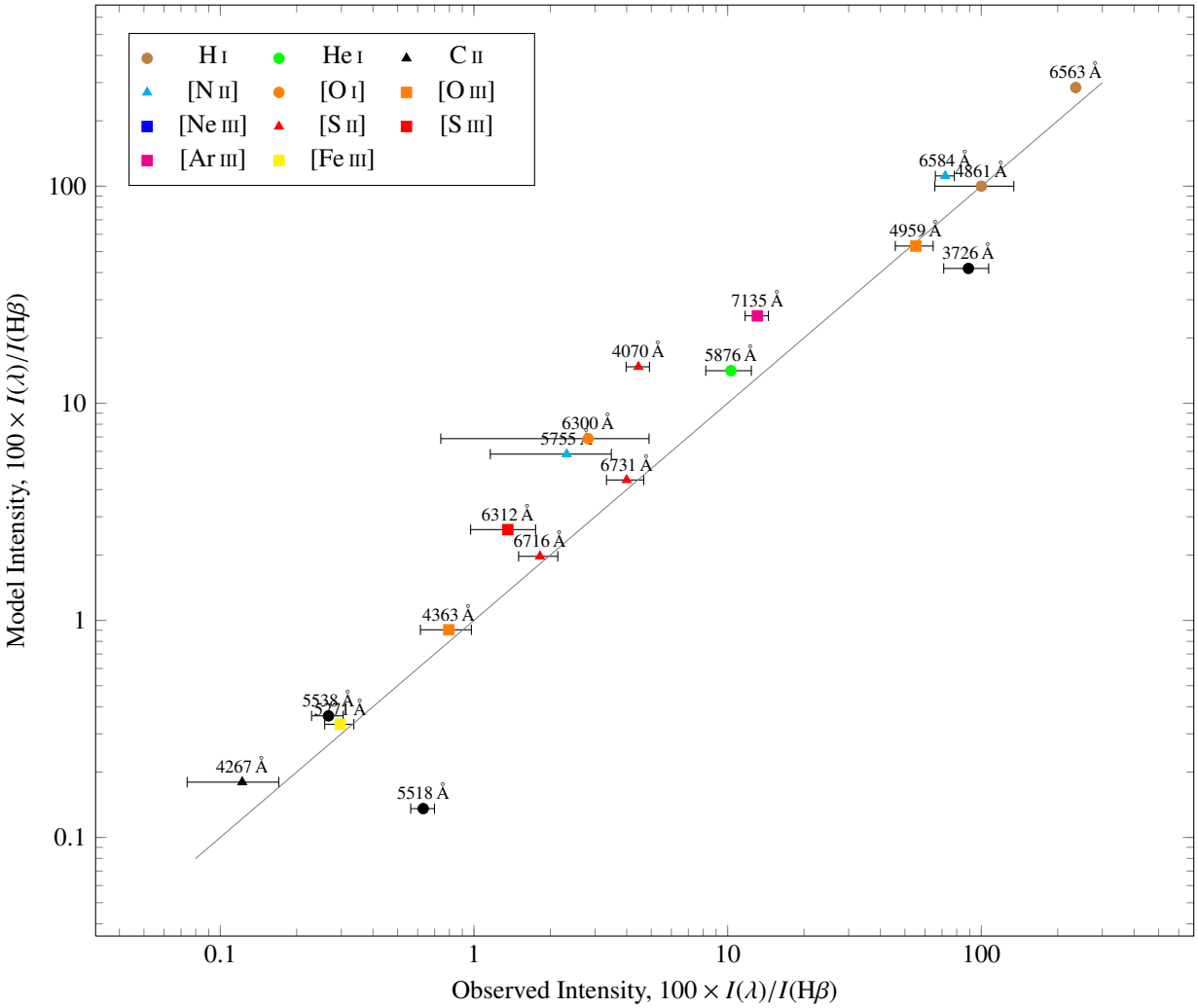


Model M: Tsamis HST10 new2

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set

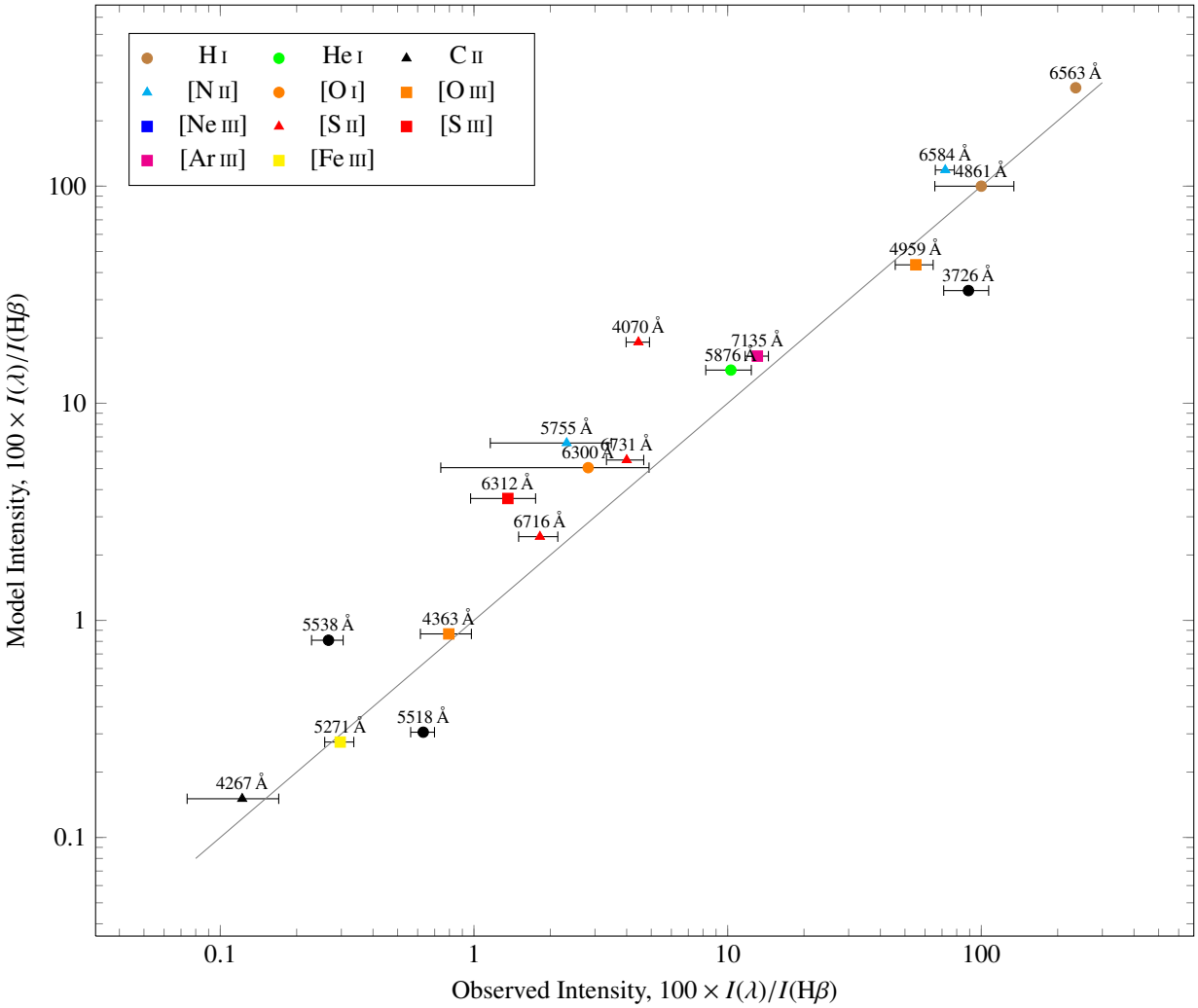


Model N: ZZ03

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.30$

Abundance set



Model O: Orion total

Spectrum WMBasic, 39 000 K

Flux $\log_{10} \Phi = 12.10$

Abundance set