



Investigating Temperature Variations of the Solar Corona during CMEs

Master's thesis submitted by

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CERTIFICATE

This is to certify that the Masters thesis project titled "**Investigating Temperature Variations of the Solar Corona during CMEs**" has been the outcome of an original study carried out by **Dheeraj Vittal Shenoy** under the supervision of **Mr. Sundar M. N.** and **Dr. Tanmoy Samanta** towards the partial fulfilment of the requirements for the degree of M.Sc. Physics of JAIN (Deemed-to-be University).

This to further certify that the work reported herein does not form a part of any other thesis/dissertation, on the basis of which a degree, diploma or a certificate has been conferred upon this or any other student in the past.

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DECLARATION

I, **Dheeraj Vittal Shenoy**, hereby declare that this dissertation titled **Investigating Temperature Variations of the Solar Corona during CMEs** has been the outcome of an original study carried out under the guidance of **Mr. Sundar M. N.** and **Dr. Tanmoy Samanta** towards the partial fulfilment of the M.Sc. Physics degree of the JAIN (Deemed-to-be University) during the year 2023-2024. This study has not been submitted for any degree, diploma or certificate.

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UNDERTAKING

I, **Dheeraj Vittal Shenoy** hereby give an undertaking that the data reported in this dissertation will not be used for any publication, conference presentation or for any industrial interaction without a written approval from the Project Supervisor and the Director, School of Sciences, JAIN (Deemed-to-be University).

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Abbreviation	Full Form
SDO	Solar Dynamics Observatory
AIA	Atmospheric Imaging Assembly
HMI	Helioseismic Magnetic Imaging
EVE	Extreme ultraviolet Variability Experiment
SOHO	Solar and Heliospheric Observatory
LASCO	Large Angle and Spectrometric Coronagraph Experiment
EUV	Extreme Ultra Violet
DEM	Differential Emission Measure
JSOC	Joint Special Operations Command
FITS	Flexible Image Transport System
SSW	Solar SoftWare
IDL	Interactive Data Language
DN	Data Number

Table 1: List of abbreviations

1 Abstract

Stellar coronal mass ejections are difficult to observe and analyse due to the lack of spatial resolution. Indirect methods have been devised to detect the signatures of CMEs on stars. We have done Sun-as-a-star analysis of CMEs for various events using the data from SDO/AIA instrument and good correlations have been observed by Differential Emission Measure (DEM) analysis.

2 Introduction

Explosive phenomena that occur on the surface of sun are called as solar flares, caused by the reconnection of a sun's magnetic field lines. These flares are often accompanied by filaments/prominent eruptions and Coronal Mass Ejections (CMEs). Flares are classified into two types: eruptive and confined. Eruptive events are the flares which are associated with CMEs and confined flares are those which are not associated with CMEs. Plasma and magnetic field structures sometimes extend outwards from the surface. These structures are called filaments/prominences. Filaments and prominences are fundamentally of the same physical properties, only difference is the angle at which it is observed. If plasma floats outside the solar limb, it is called prominence, it is called a filament if it is within the solar background otherwise. Difference is in the type of spectrum obtained, which is Balmer lines in emission spectra in case of prominences and absorption lines in case of filaments. Loop-like structures of plasma stand out brightly against the dark background of space, these are prominences. Some prominences appear dark compared to the bright background of the Sun, these are called filaments.

CMEs are large plasma structures ejected from the solar surface to the heliosphere. They were first discovered in 1971 ([Gopalswamy, 2016](#)). They majorly affect the space weather and are sometimes dangerous when their trajectory is towards Earth. Similar to solar flares, stellar flares maybe associated with CMEs, but it is not easy to detect or study them as there is no spatial resolution unlike Sun. CMEs accompanied by stellar flares are substantially larger in comparison to the ones observed on the Sun, and have been known to affect the exoplanets around the host stars. Hence studying the CMEs is very crucial for future exoplanetary expeditions also.

Indirect stellar CME detections have been explored with methods such as coronal dimming, blueshifted emission of chromospheric lines, X-ray, EUV and FUV dimming, Type-II and type-IV radio bursts etc. Among this, coronal dimming is considered one of the promis-

ing signatures to detect the occurrence of stellar CMEs (**Namekata et al., 2022**). Regions with temporary dimming of plasma on the solar surface is seen after an eruptive event like CMEs in the EUV and soft X-ray wavelengths. Dimming is most prominently observed in the 1-2 million K range of Sun's plasma of the quiet corona. Time series plot of irradiance value of Sun shows a prominent decrease after the event compared to the value before the event. This effect is known as Coronal Dimming. The probability of Coronal Dimming being associated with CMEs have been observed to be very high $P(Dim | CME) = 0.842$ in comparison to false alerts $P(Dim | !CME) = 0.167$ (**Veronig et al., 2021**). Many studies have been done to get information about the underlying CME from the coronal dimming, like to get the information about the mass and velocity of the CME uamount of depth and slope of the curve of the CME have been analysed to get information about their mass and velocity (**Mason et al., 2016**).

In this study, we investigate the temperature distribution of the corona during CMEs and formulate a new method for Sun-as-a-star analysis of CMEs with the help of Differential Emission Measure (DEM) analysis.

2.1 Instrument

We are using Atmosphere Imaging Assembly (AIA) instrument of Solar Dynamics Observatory (SDO; (**Pesnell et al., 2011**)) for the Sun's spectral irradiance data. SDO is a space observatory launched by NASA on 2010 as a part of ‘Living With a Star’ (LWS) program. The spacecraft contains three instruments on board: Extreme Ultraviolet Variability Experiment (EVE), Helioseismic and Magnetic Imager (HMI), Atmospheric Imaging Assembly (AIA). We'll be focusing on the AIA instrument, since that is what we are using. AIA was built in partnership with Lockheed Martin Solar and Astrophysics Laboratory (LMSAL). AIA contains 4 cassegrain telescopes which are optimized to observe narrow bands in the EUV region. Each of the four f/20 telescope has a 20-cm primary mirror and an active secondary mirror. The telescope is designed to prevent charged particles from reaching the

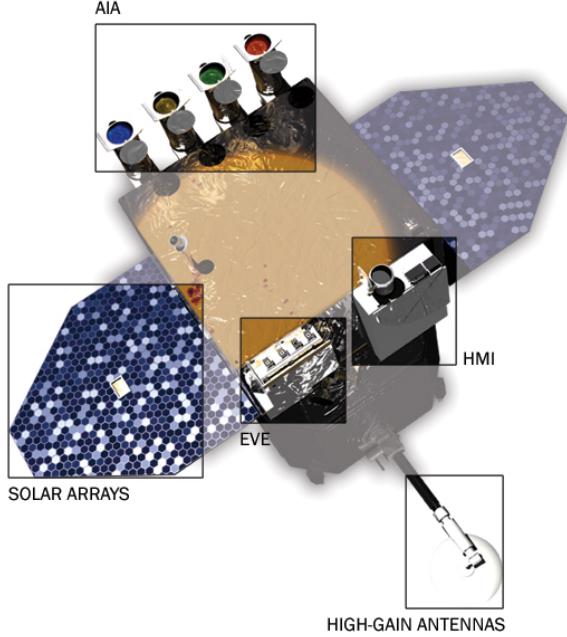


Figure 1: Solar Dynamics Observatory Spacecraft. Image obtained from <https://sdo.gsfc.nasa.gov/mission/spacecraft.php>

Charge Coupled Device (CCD). Field of View (FOV) of each of the telescope is 41 arcmin in circular diameter. The mirrors have special multilayer coatings that are optimized to observe the selected EUV wavelengths of interest. Three of the telescopes have two different EUV bandpasses. The CCDs are back-thinned and back-illuminated with 4096×4096 pixels capturing capabilities. Each of the $12 \mu\text{m}$ pixel corresponds to 0.6 arcsec. The telescopes have a selector mechanism to choose the wavelength. AIA captures full-frame EUV image and one UV or visible-light image every 12 seconds. (**Lemen et al., 2011**)

The different channels of AIA respond differently to the radiation of different temperature. Temperature response curves gives the information about the response of each of the channels with respect to the temperature of the radiation being received. Figure 4 shows the response curves, which has been obtained using `aia_get_response.pro` procedure of SSW IDL.

Band	Primary role, ion(s)	Region of Sun's atmosphere	logT[K]
4500 Å	Continuum	Photosphere	3.7
1700 Å	Continuum	Temperature minimum, photosphere	3.7
304 Å	He II	Chromosphere, transition region	4.7
1600 Å	C IV, continuum	Transition region, upper photosphere	5.0
171 Å	Fe IX	Quiet corona, upper transition region	5.8
193 Å	Fe XII, XXIV	Corona and hot flare plasma	6.1, 7.3
211 Å	Fe XIV	Active region corona	6.3
335 Å	Fe XVI	Active region corona	6.4
94 Å	Fe XVIII	Flaring regions	6.8
131 Å	Fe XX, XXIII	Flaring regions	7.0, 7.2

Table 2: AIA wavelength channels. Table obtained from <https://aia.lmsal.com/public/instrument.htm>



Figure 2: SDO AIA Telescope (Lemen et al., 2011)

3 Methodology

In this study, we wish to find a method of detecting signatures of CMEs on the Sun by converting it to a point source (hereafter referred to as **Pointifying**) and then inspecting if the signatures of the CMEs exist after the conversion and is similar to what was before the conversion. Pointifying the Sun roughly translates to converting Sun to a star, or placing Sun to a place that's distant from Earth/observer in comparison to the distance between us

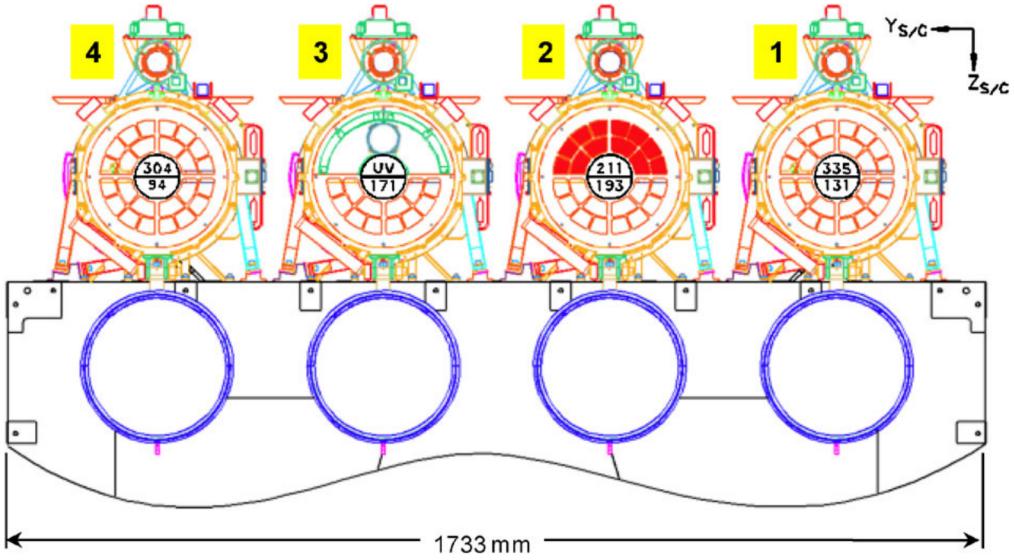


Figure 3: AIA Telescope Layout. Telescope 2 has an aperture blade to select between it's two wavelength channels. Rest of the telescopes rely on the filters in filter wheels to select the channel ([Lemen et al., 2011](#))

and the Sun (defined as astronomical unit ($1 \text{ AU} \approx 1.496 \times 10^8 \text{ km}$)) such that it appears as a point source. We then analyse and compare the irradiance from the point source Sun and the full disk Sun using DEM to see if they show similar signatures of CMEs. Typically, Sun-as-a-star analysis involves selecting a region of interest on the surface of the Sun, making the assumption that this is the only region that affects the event under study, and that there is no activity anywhere on the rest of the Sun's surface, and then integrating the parameter of interest over the entire Solar disk. This is again a rough approximation to an actual star.

In the following section, we discuss about the event selection, data used for the study, data analysis and finally the results and conclusion.

3.1 Data Analysis tools used

We have made use of Python programming language for our analysis. We have made extensive use of the following libraries: Numpy, Scipy, Aiapy, Sunpy, Pandas, Matplotlib, Astropy, Natsort, Multiprocessing, Datetime, Movieipy. In addition, we have used the RML method code for DEM analysis, as mentioned in ([Massa et al., 2023](#)). We have made use

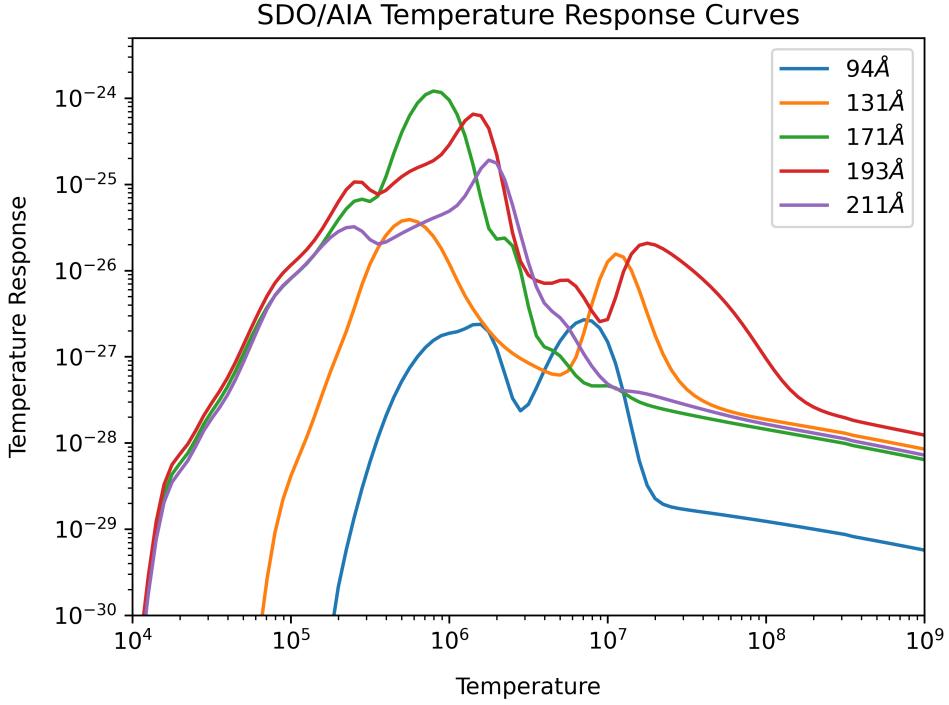


Figure 4: SDO/AIA Temperature Response Curves obtained using `aia_get_response.pro`. Curves corresponding to the five wavelength channels required for the DEM analysis has been plotted. Temperature has usual unit of [K] and temperature response function has the unit of [$\text{DN cm}^{-5} \text{s}^{-1} \text{pixel}^{-1}$]

of SSW IDL for obtaining the temperature response curves for AIA. For visual inspection of the CME events, we have used JHelioViewer software. For quick inspection of the data, we have used FITSExplorer¹, a software created as a side project by D. V. Shenoy.

3.2 Event Selection

We have chosen three CME events that have erupted on 2011 August 04, 2012 August 31 and 2021 October 28. The Solar and Heliospheric Observatory's (SOHO) Large Angle and Spectrometric Coronagraph Experiment (LASCO) C2 (SOHO/LASCO) images of the CMEs obtained from the SOHO/LASCO CME catalog (https://cdaw.gsfc.nasa.gov/CME_list/) are given in fig. 6 and fig. 7. Brief description of the events is given below:

¹link to the github repo: <https://github.com/dheerajshenoy/FITSExplorer>

1. **2011 August 04:** This event has been referred from ([Mason et al., 2016](#)) in which it is the 20th Event. The event started at around 04:12 UT. This event occurred from the source location N19W36 associated with the active region AR 11261.
2. **2012 August 31:** This CME event was associated with a long filament eruption and it erupted around 19:49 UT. The CME associated with the filament travelled at over 900 miles per second.

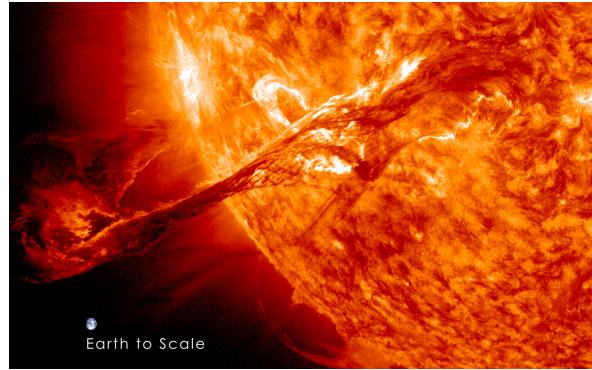


Figure 5: Image of the Earth to scale with the filament eruption of 31st August 2012. Image credit: <https://svs.gsfc.nasa.gov/11095>

3. **2021 October 28:** This is an example for rarely occurring ‘ground level enhancement’ event. During such an event, particles from the Sun are energetic enough to pass through the magnetic sheath that surrounds Earth and protects us from low energy solar outbursts. This was only the 73rd ground level enhancement since records began in the 1940s, and none have been recorded since ([Klein et al., 2022](#)). The event occurred around 15:17 UT. The flare associated with the CME was an X1.1 class flare originated from the active region AR 2887.

We have used 10 hours of data for the first two events, and about 7 hours of data for the third event. All three event data are at 2 minute cadence.

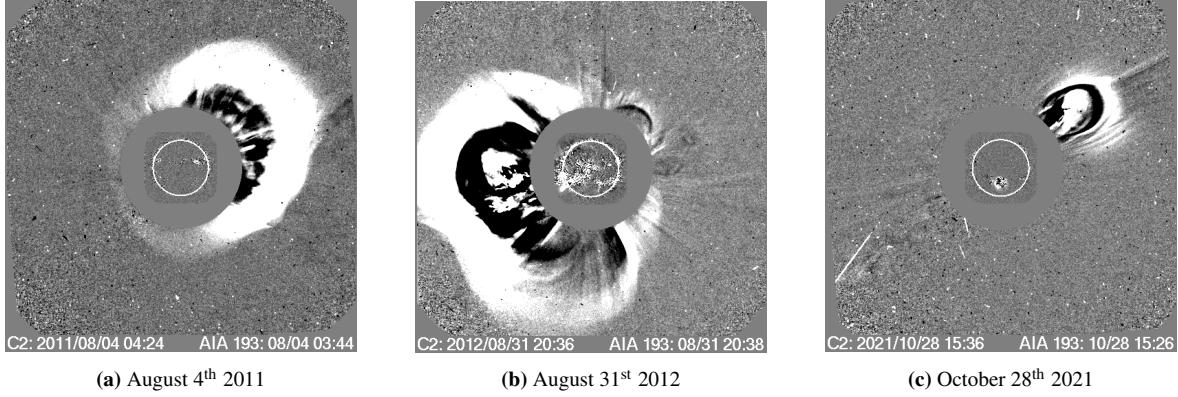


Figure 6: SOHO/LASCO C2 images of the three selected events

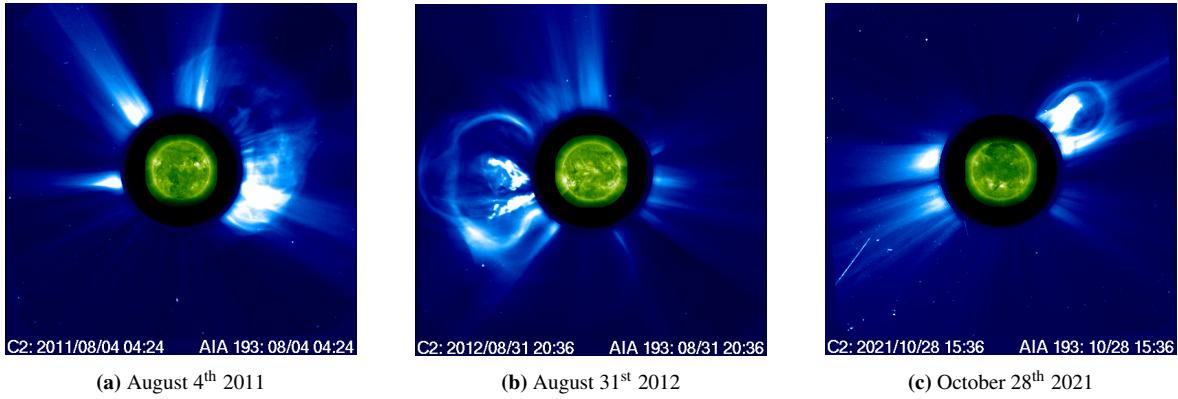


Figure 7: SOHO/LASCO C2 AIA 193 Åfilter images of the three selected events

3.3 Data

The SDO/AIA data is accessed through the JSOC portal (<http://jsoc.stanford.edu>) and required event data is obtained through the service. Event data consists of FITS files of Sun's image for the selected wavelength bands. For our analysis, we have used 5 channels: 94 Å, 131 Å, 171 Å, 193 Å and 211 Å. The remaining channels probe the Sun's surface temperature that is greater than what is required for our analysis. Also, the 335 Å channel has been excluded because of it's relatively weak temperature response at any temperature which affects the RML method used for the DEM analysis (**Massa et al., 2023**).

3.4 Image Pre-Processing

The downloaded FITS data files are 4096×4096 pixels in dimension. Data downloaded from the portal is level 1, which has been flat-fielded and processed to remove bad pixels and spikes (only for EUV channels), but not registered to preserve precise pixel values. As different channels of AIA have different roll angles, multi-wavelength analysis of any kind with level 1.0 data is problematic. Also, the pointing information contained in the headers of these FITS images will not be accurate, as it would have undergone changes compared to the information stored when the image was created. As mentioned in the SDO Data Analysis Guide (REFERENCE REQUIRED), we have to use `aia_prep.pro` function in Solar Software (SSW) IDL to correct or prepare the data used. The images are downscaled from their original 4096×4096 pixel dimension to 512×512 pixels using sunpy as obtaining DEM solutions for the 4k dimension would be really time consuming and unnecessary as we are comparing full disk and point source DEM solutions.

The **aiapy** library is used to carry out the necessary procedure like ‘**Pointing correction**’ and ‘**Registration**’ as mentioned in the documentation of aiapy, to convert level 1 data to level 1.5. Pointing correction updates the keywords in the header of the FITS file to the latest information and Registration rotates, scales and translates the image so that the Solar north is aligned with the solar north and each pixel is 0.6 arcsec cross, and the center of the Sun is at the center of the image. Now, after this calibration, the images are good for multi-wavelength analysis. Finally, the images are normalized with respect to their exposure time. This is done because the images are captured under different lighting conditions or exposure settings, which leads to incorrect analysis when doing a multi-wavelength comparison. Without this, differences in brightness due to varying exposure times could distort interpretations and analysis.

3.5 DEM Analysis

Direct analysis of light curve might seem like a good choice as the effects of CMEs are seen in light curves too. But, light curve doesn't have any information about the temperature of the plasma that is being expelled. Information regarding the temperature can be derived using Differential Emission Measure (DEM) solutions obtained using the images of Sun.

3.5.1 Emission Measure

Emission measure (EM) is a quantity used in astrophysics to describe the amount of emitting material along the line of sight in a particular volume, usually in the context of a hot or ionized gas, such as a stellar atmosphere. It provides a measure of the emission intensity of a given region at various temperatures. Emission measure is expressed in units of cm^{-3} or cm^{-5} , representing the number of particles emitting radiation per unit volume or per unit area, respectively.

Emission measure is related to the number density of particles n in a volume dV of plasma, in a particular temperature range T_1 and T_2 , along the line of sight, is given by,

$$EM = \int_{T_1}^{T_2} n^2 dV$$

Emission measure is a crucial parameter in understanding the energetics and physical conditions of a plasma, such as those found in stars, galaxies and other astrophysical environments. In the context of the Sun, the emission measure is often used to study the solar corona, helping us to understand the distribution of temperatures and the processes governing the heating of the outer solar atmosphere.

3.5.2 Differential Emission Measure

Differential Emission Measure (DEM) is used to describe the distribution of emitting material at different temperatures in a given volume. It is a measurement of the amount of plasma

at various temperatures per unit volume along the sight. DEM helps us understand how much material is present at different temperature in stellar atmosphere. This is crucial for studying the physical conditions and processes occurring in stellar atmospheres. DEM is usually expressed in units of $cm^{-5}K^{-1}$, representing the number of particles emitting radiation at a particular temperature per unit volume.

$$DEM = f(T) = \frac{d}{dT} EM = n^2 \frac{dV}{dT}$$

The integral of $DEM(T)$ over a finite temperature range is called as the emission measure. This quantity helps to understand the thermal structure of a stellar atmosphere, providing insight into the distribution of temperatures and the heating mechanisms that operate in a particular region. DEM arises from certain aspects of coronal emission line. Optically thin property of corona, scaling of emission line intensity with density squared n^2 (for most lines) and temperature response function, $R(T)$, that peaks at certain temperature for each of the lines.

If $R(x) \equiv R(T(x))$ is the temperature response function, then, line intensity along the line of sight can be written as,

$$I \propto \int_{LOS} n^2(x)R(T(x)) dx$$

From a DEM , parameters like plasma density, thermal X-ray flux, thermal energy and weighted temperature emission measure etc. can be estimated (**Su et al., 2018**). In solar research, DEM aids in understanding the Sun's atmosphere, while in stellar astrophysics, it contributes to characterizing other stars, enhancing our knowledge of stellar diversity and evolution (**Namekata et al., 2023**).

3.5.3 Differential Emission Measure Inversion

DEM inversion refers to the process of determining the physical conditions of the plasma from observed coronal emission line data. Radiation emitted by the plasma is observed at different wavelengths using spectroscopic techniques. The observed data is then used to construct the DEM, which represents the distribution of emitting material at different temperatures in the stellar atmosphere. Next, the inversion process is employed, which is basically reconstructing the DEM, thereby helping us to determine the underlying temperature distribution which gave rise to the observed emission line intensities. Then different computational techniques can be employed to fit theoretical models of the emission at different temperatures to the observed data. The best fit model then provides the information about the temperature distribution of the plasma.

Many DEM inversion techniques have been developed over the years: basis pursuit technique ([Cheung et al., 2015](#)), fast iterative regularized method ([Plowman et al., 2013](#)), iterative SITES method ([Morgan and Pickering, 2019](#)), regularized method (REG) ([Hannah and Kontar, 2012](#)), regularized maximum likelihood (RML) method ([Massa et al., 2023](#)) etc. We will be making use of the RML method, as it has been found to be a good approximation to the actual DEM profiles and is performant in comparison to the other methods.

3.6 Data Analysis Procedure

After calibration (registering, pointing correction and exposure time normalization), the image data is fed to the RML code ([2023](#)) which returns the DEM solutions for the desired temperature range. We choose a temperature (logarithm) range of $\log T[K] = [5.85, 6.4]$ which translates to temperature of $10^{5.85} \approx 7 \times 10^5 K$ to $10^{6.75} \approx 5.6 \times 10^6 K$. The function `rml_dem` takes the following inputs: array containing the AIA data (DN/s), array of uncertainties of AIA data (DN/s), exposure time value for each AIA channel (s), array containing the temperature response function for each channel ($DN \text{ cm}^5 \text{ pixel}^{-1} \text{ s}^{-1}$), array containing the value of log base 10 of the center of each temperature bin, array containing the value of

the width of base 10 logarithm of the temperature bins. The function returns an array containing the values of the reconstructed DEM profiles. The returned solutions will be spaced according the temperature spacing of the response function array fed to the solver.

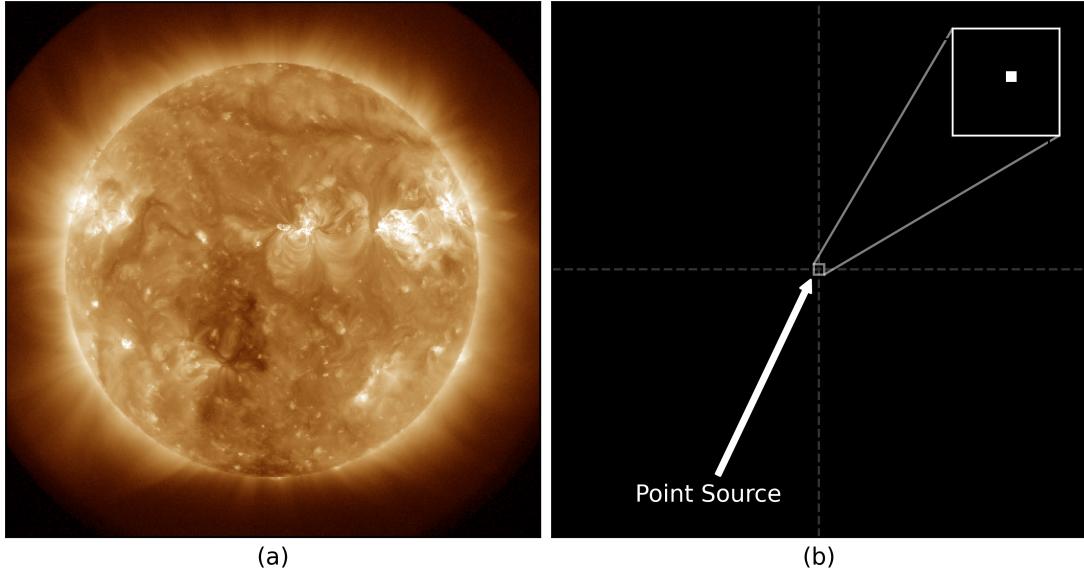


Figure 8: (a) shows the full disk image of Sun in the 193 Å channel, (b) shows the image after pointification.

Next step is to generate the point source from the full disk, which we refer to as “Pointification”. This involves the conversion of the full disk image of Sun to a point source, which mimics a distant star. The original image of 4096×4096 pixels dimension is averaged and new image having 512×512 pixels dimension is created with the pixel value at (256, 256) of the image (midpoint of the image) equal to the calculated average image pixel value. The average values of 4096×4096 pixel image dimension has been considered and not the 512 one so as to not allow the errors due to the resampling of the image affect the average pixel intensity value significantly. Figure 8 shows the result of pointification procedure. After pointification, DEM profile is reconstructed from these images. These act like DEM profiles of a point source/star.

4 Results

Figure 9, fig. 10 and fig. 11 shows the timeseries plot of DEM for the three events. The time series plot is calculated by averaging the acceptable DEM solutions obtained for each image over three consecutive temperature range (i.e averaging the solutions for 5.85, 5.9 and 5.95 etc.). Blue curve corresponds to the full disk DEM and red curve corresponds to the point source DEM. The temperature range above $\log T = 6.45$ has been omitted as no correlation was found between point source and full disk.

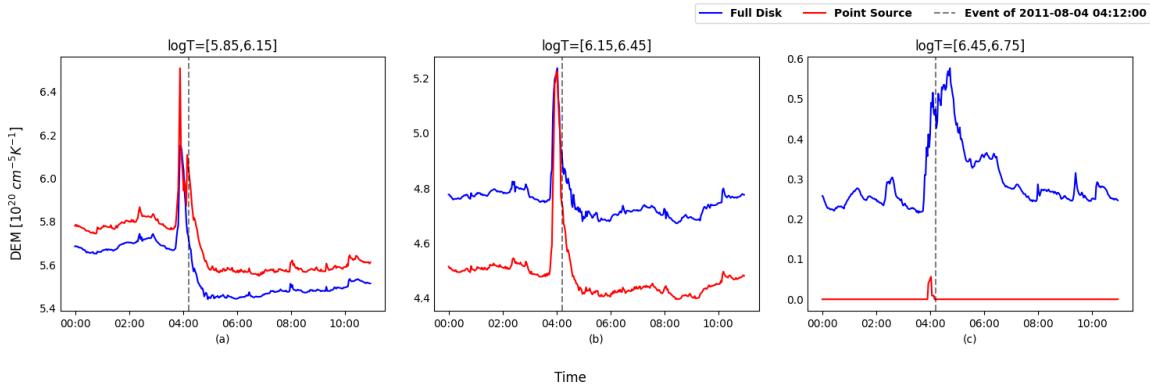


Figure 9: Timeseries of DEM for 4th August 2011 Event.

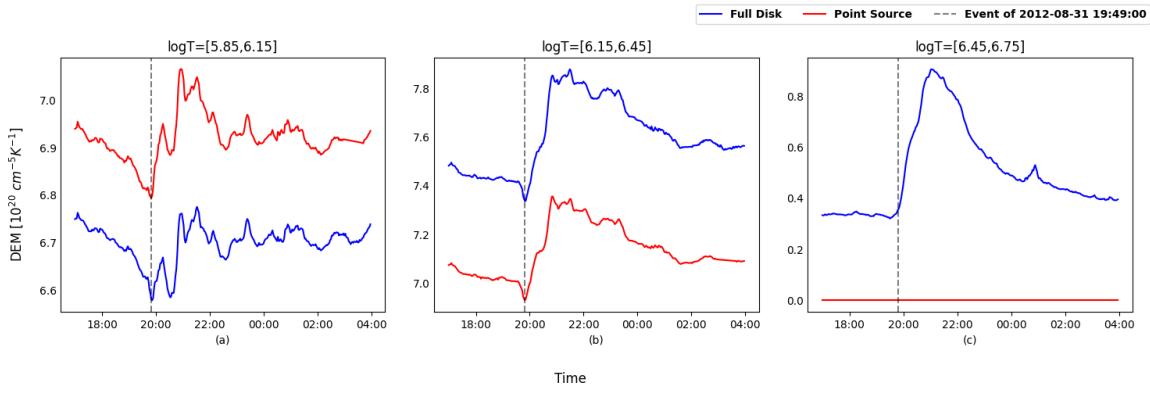


Figure 10: Timeseries of DEM for 31st August 2012 Event

The sudden increase in the DEM curve is due to the solar flare which is associated with the CME. From the above figures, we can see that there is very high correlation for the first two temperature ranges, but there is little to no correlation in the DEM profiles of the full

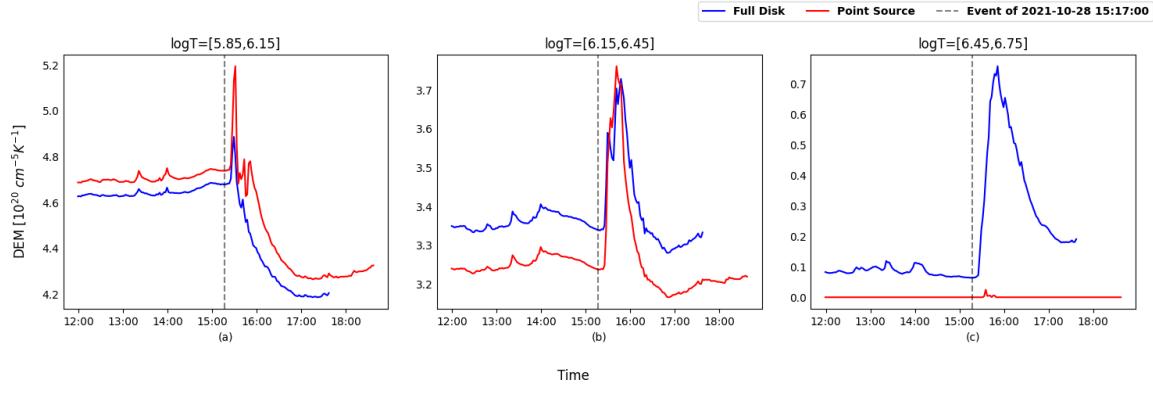


Figure 11: Timeseries of DEM for 28th October 2021 Event

Event	Pearson Correlation Coeffecient		
	logT=[5.85, 6.15]	logT=[6.15, 6.45]	logT=[6.45, 6.75]
August 2011	0.9449	0.9767	0.2190
August 2012			
October 2021	0.9555	0.9577	0.2578

Table 3: Correlation between Point source and Full Disk DEM

disk and point source for the temperature range $\text{logT}=[6.45, 6.75]$. For temperature ranges greater than $\text{logT}=6.45$, the correlation is almost 0. We make use of Pearson's Correlation coeffecient to find out the amount of correlation between the point source and full disk DEM, for which we use `pearsonr` function from the `scipy` library in Python.

We see a discrepancy in the value of DEM between the point source and full disk average values. This could be due to the error induced during the DEM profile reconstruction, instrumental errors, error incurred during the resampling or reduction of image dimension from 4096×4096 pixels to 512×512 and it could also be due to averaging error. The dataset length is not equal in some of the case as invalid or small DEM solution values have been removed.

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