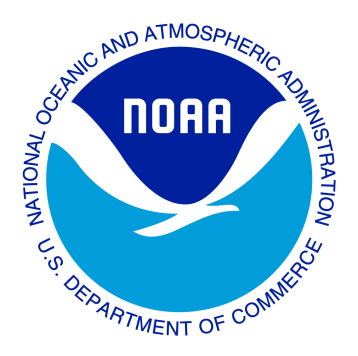
Climate Data Record (CDR) Program

Climate Algorithm Theoretical Basis Document (C-ATBD)

Solar Irradiance



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TABLE of CONTENTS

1. Introduction 1

1.1 Purpose 1

1.2 Definitions 1

1.3 Document Maintenance 2

2. Overview of Solar Irradiance Climate Data Record 3

2.1 Products Generated 3

2.2 Instrument and Model Characteristics 4

3. Algorithm Description 5

3.1 Algorithm Overview 5

3.2 Processing Outline 5

3.3 Algorithm Input 7

3.3.1 Primary Input Data 7

3.3.2 Ancillary Data 8

3.3.3 Derived Data 9

3.3.4 Forward Models 10

3.4 Theoretical Description 10

3.4.1 Physical and Mathematical Description 15

3.4.2 Data Merging Strategy 18

3.4.3 Numerical Strategy 18

3.4.4 Calculations 18

3.4.5 Look-Up Table Description 18

3.4.6 Parameterization 18

3.4.7 Algorithm Output 18

4. Test Datasets and Outputs 22

4.1 Test Input Datasets 22

4.2 Test Output Analysis 24

4.2.1 Reproducibility 24

4.2.2 Precision and Accuracy 24

4.2.3 Error Budget 24

5. Practical Considerations 26

5.1 Numerical Computation Considerations 26

5.2 Programming and Procedural Considerations 26

5.3 Quality Assessment and Diagnostics 26

5.4 Exception Handling 27

5.5 Algorithm Validation 27

5.6 Processing Environment and Resources 29

6. Assumptions and Limitations 30

6.1 Algorithm Performance 33

6.2 Sensor Performance 33

7. Future Enhancements 36

7.1 Enhancement 1 36

7.2 Enhancement 2 37

7.3 Enhancement 3 37

8. References 39

Appendix A. Acronyms and Abbreviations 41

LIST of FIGURES

Figure 1: Flow diagram of the Algorithm Processing

Figure 2: Facular brightening and sunspot darkening time series during the solar cycle.

Figure 3: Facular brightening and sunspot darkening time series during solar rotation.

Figure 4: Reference Spectrum of the quiet Sun in NRLSSI2.

Figure 5: Scaling coefficients that convert the facular brightening and sunspot darkening indices in Figures 2 and 3 to their equivalent irradiance change, in energy units.

Figure 6: Total solar irradiance reconstructed with the NRLTSI2 model (upper), and the corresponding the facular and sunspot contributions (lower).

Figure 7: Spectral irradiance in broad wavelength bands reconstructed with the NRLSSI2 model.

Figure 8: Comparison of NRLTSI2 total irradiance calculated by the C-ATBD algorithm, with the integral of the NRLSSI2 spectra also calculated by the algorithm, upper, and of their respective facular and sunspot components (lower).

Figure 9: Comparison of NRLTSI2 total solar irradiance, calculated by the C-ATBD algorithm with an earlier model, NRLTSI.

Figure 10: Figure 10: Comparison of NRLSSI2 solar spectral irradiance in the selected broad wavelength bands shown in Figure 7, calculated by the C-ATBD algorithm, with an earlier model, NRLSSI. On the left are the time series and on the right are the percentage differences.

Figure 11: Comparison modeled total solar irradiance time series with observations made by TIM on SORCE, during solar rotation (upper panel) and during the solar cycle (middle and lower panels).

Figure 12: Ratios of NRLTSI2 and NRLTSI models to TIM observations of total solar irradiance variability.

Figure 13: Comparisons of NRLSSI2 solar spectral irradiance in broad wavelength bands with SIM observations. Both the model and observed time series have been detrended; the variations are associated with the Sun’s 27-day rotation.

Figure 14: Solar cycle (24) changes shown as percentages (upper) and in energy units (lower) in the NRLSSI2 model, compared with NRLSSI.

LIST of TABLES

Table 1: Definitions of symbols used in the C-ATBD.

Table 2: Products that this CDR provides.

Table 3: Structure of NRLTSI2 output for daily values of TSI from 1978 to the present.

Table 4: Structure of NRLSSI2 output for daily values of SSI from 1978 to the present.

Table 5: Estimated uncertainties for the solar irradiance calculated by the algorithm.

Table 6: Summary of assumptions in the theoretical basis for modeled solar irradiance, model inputs and the potential validation approaches.

1. Introduction
   1. Purpose

The purpose of this document is to describe the algorithm submitted to the National Climatic Data Center (NCDC) by Judith Lean (Space Science Division, Naval Research Laboratory), Odele Coddington and Peter Pilewskie (Laboratory for Atmospheric and Space Science, University of Colorado) that is used to create the Solar Irradiance Climate Data Record (CDR). Also described are the solar activity indices of sunspot darkening, S(t) ,and facular brightening, F(t), which are input to the algorithm. The algorithm calculations of solar irradiance augment direct measurements made by the Total and Spectral Solar Irradiance Sensor (TSIS). The actual algorithm is defined by the computer program (code) that accompanies this document; this C-ATBD provides a guide to understanding that algorithm, from both a scientific perspective and in order to assist a software engineer or end-user performing an evaluation of the code.

This C-ATBD for the Solar Irradiance Climate Data Record describes the procedures, algorithms and input datasets used to construct the total solar irradiance and the concurrent spectral solar irradiance variations during recent decades and in historical time periods since 1615. Also described are the output files of daily, monthly and annual solar irradiance, and validation procedures for the modeled solar irradiance. The solar irradiance reconstructions that this C-ATBD describes compliment the direct measurements made of total and spectral solar irradiance by the TSIS instrument, documented in the TSIS ATBD.

* 1. Definitions

**Table 1. Definitions of symbols used in the C-ATBD**

|  |  |  |
| --- | --- | --- |
| **Symbol** | **Definition** | **Units** |
| t | time |  |
| λ | wavelength | nm |
| T(t) | total solar irradiance at time t | W m-2 |
| TQ | total solar irradiance (TSI) of the quiet (inactive) Sun; invariant with time | W m-2 |
| I(λ,t) | solar spectral irradiance (SSI) at wavelength λ and time t | W m-2 per wavelength bin |
| IQ(λ) | solar spectral irradiance of the quiet (inactive) Sun at wavelength λ | W m-2 per wavelength bin |
| F(t) | facular brightening index at time t | units of the GOME Mg index |
| S(t) | sunspot darkening index at time t | millionths of solar hemisphere |
| FQ | facular brightening of the quiet (invariant) Sun, corresponding to IQ and SQ | units of the GOME Mg index |
| SQ | Sunspot darkening of the quiet (invariant) Sun, corresponding to IQ and FQ | millionths of solar hemisphere |
| AS | sunspot area | millionths of solar hemisphere |
| μ | sunspot location in radial heliocentric coordinates |  |
| Mg(t) | The Mg II index at time t, determined as the ratio of core to wing emission in the Mg II Fraunhofer line | units of the GOME Mg index |

* 1. Document Maintenance

The algorithm that calculates the total solar irradiance (TSI) and solar spectral irradiance (SSI) that compose the Solar Irradiance Climate Data Record uses bolometric (for TSI) and wavelength-dependent (for SSI) coefficients to convert two time-varying inputs, the facular brightening index, F(t), and the sunspot darkening index, S(t), into their respective modulations of the solar irradiance of the “quiet” Sun (i.e., in the absence of these activity features). The starting point for the algorithm is a baseline (invariant) value of total, TQ, and spectral, IQ(λ), irradiance representative of the “quiet” (inactive) Sun, obtained directly from space-based irradiance observations. The numerical values of the coefficients that determine the irradiance modulation imposed by sunspots and faculae, when these features are present on the solar disk, are determined from analysis of variability in extant TSI and SSI datasets measured by instruments onboard the Solar Radiation and Climate (SORCE) spacecraft (Rottman et al., 2005). Neither the baseline irradiance values nor the coefficients that calculate the irradiance modulation with time are expected to change on time scales shorter than a few years, which is the time frame over which new information can be expected from ongoing TSIS measurements. For this reason, it is not expected that the algorithm will evolve rapidly, nor that frequent synchronization will be needed.

1. Overview of Solar Irradiance Climate Data Record
   1. Products Generated

The Solar Irradiance Climate Data Record provides values of the total and spectral irradiance as functions of time, listed in Table 1. Total solar irradiance, T(t), is the total, spectrally integrated (i.e., bolometric) energy input to the top of the Earth’s atmosphere, at a standard (invariant) distance of one Astronomical Unit (1AU) from the Sun. Its units are W m-2. Solar Spectral Irradiance, I(λ,t), is the corresponding spectrum that integrates self-consistently to the total. Values of spectral irradiance are provided in 3785 (variable width) wavelength bands from 115.5 ± 0.5 nm to 99,975 ± 25 nm, in units of W m-2 per band.

Both direct observations and model calculations contribute to the Solar Irradiance Climate Data record. The Total and Spectral Irradiance Sensor (TSIS) is the corresponding observing system that measures the total solar irradiance and solar spectral irradiance directly. It is described in a separated ATBD (**ref**). The Naval Research Laboratory Total Solar irradiance (NRLTSI2) and Naval Research Laboratory Solar Spectral irradiance (NRLSSI2) models, developed from the SORCE database of solar irradiance observations, calculate the irradiance from two primary (time varying) inputs, the sunspot darkening, S(t), and facular brightening, F(t).

Also provided are four reference spectra indicative of quiet, moderate and high solar activity levels, and the Maunder Minimum, in 99844 bins of equal 1 nm width, from 115.5 to 99999.5 nm.

Table 2: Products that this CDR provides.

|  |  |  |  |  |
| --- | --- | --- | --- | --- |
| Product | Type | Number of Wavelength Bins | Time Range | Cadence |
| Total Solar Irradiance | Observational composite | 1 | 1978-presemt | daily |
| Total Solar Irradiance | NRLTSI2 model | 1 | 1882-present | daily, monthly |
| Total Solar Irradiance | NRLTSI2 model | 1 | 1610-present | annually |
| Solar Spectral Irradiance | NRLSSI2 model | 3785  (variable width) | 1882-present | daily, monthly |
| Solar Spectral Irradiance | NRLSSI2 model | 3785  (variable width) | 1610-present | annually |
| Solar Spectral Irradiance | Reference Spectrum | 99884  (1 nm width) | Quiet, moderate & high solar activity | NA |
| Solar Spectral Irradiance | NRLSSI2 model spectrum | 99884  (1 nm width) | Maunder Minimum  (estimated) | NA |

* 1. Instrument and Model Characteristics

The primary solar irradiance observational dataset for the Solar Irradiance Climate Data Record is that measured by the Total Solar and Spectral Irradiance Sensor (TSIS), which a separate ATDB document describes (**ref**).

The present document describes the algorithm that calculates solar total and spectral irradiance using the NRLTSI2 and NRLSSI2 models, developed from direct solar irradiance observations made by Total Irradiance Monitor (TIM) and Solar Irradiance Monitor (SIM) instruments on the Solar Radiation and Climate Experiment (SORCE) spacecraft from 2003 to 2014. The NRLTSI2 and NRLSSI2 models are the second generation of the NRLTSI and NRLSSI solar irradiance variability models that input proxy indicators of faculae and sunspots to calculate the change in a reference spectrum that these features produce, when present on the solar disc. Lean and Woods (2010) provide an overview of the NRLTSI and NRLSSI models.

The total solar irradiance generated by the NRLTSI2 model is directly comparable with daily average values of the total solar irradiance measured by the next-generation TIM, which is part of TSIS. The solar spectral irradiance generated by the NRLSSI2 model is directly comparable with daily average values of the next-generation SIM, which is also part of TSIS, after adjustment to match the wavelength bins on which the TSIS SIM measurements are reported.

1. Algorithm Description
   1. Algorithm Overview

The Solar Irradiance Data Record algorithm uses the NRLTSI2 and NRLSSI2 models to calculate values of the total irradiance and the solar spectral irradiance in 3,785 (variable) wavelength bins (on a wavelength grid from 115.5 to 99975 nm) when supplied with two inputs; the facular brightening, F(t), and sunspot darkening, S(t), indices, each of which varies with time, t.

The facular and sunspot influences on solar irradiance are calculated in Wm-2 by applying (constant) scaling coefficients to the input facular and sunspot indices, and the resultant irradiance increments are then applied to adjust the specified (invariant) baseline total irradiance and spectral irradiance which indicate the “quiet” sun. For the purposes of the algorithm, the quiet sun irradiance is that corresponding to the absence of both sunspots and faculae. An additional third component is calculated to estimate the contribution of an assumed long-term facular component speculated to produce secular irradiance change underlying the solar activity cycle on historical time scales (prior to 1950), since 1615, including during the Maunder Minimum of anomalously low solar activity.

Table 2 summarizes the outputs of the algorithm.

* 1. Processing Outline

The flow diagram in Figure 1 provides an overview of algorithm processing steps in calculating the total solar irradiance, T(t), and solar spectral irradiance, I(λ,t), at a specified time, t. The algorithms assumes that when faculae and sunspots are present on the solar disc, they alter the baseline (quiet) total solar irradiance, TQ, by amounts and , respectively:

Similarly the faculae and sunspots alter the baseline solar spectral irradiance, IQ(λ), by wavelength-dependent amounts, and :

such that the integrated spectral irradiance equals the corresponding total irradiance:

The first step in the algorithm is determination of the corresponding values for the sunspot darkening, S(t), and facular brightening, F(t), indices at time, t, which produce incremental changes in total solar irradiance:

and in solar spectral irradiance:

such that

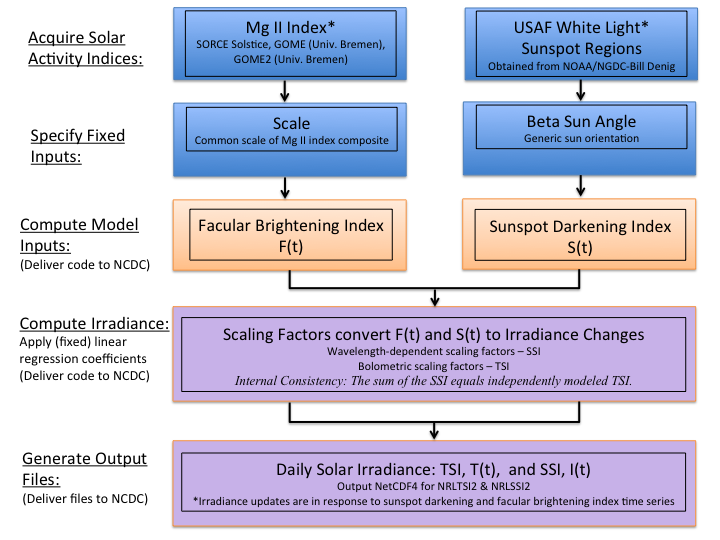
The facular brightening index is obtained by downloading the space-based Mg II composite index of solar activity constructed by combining observations made by the GOME, SCIAMACHY and GOME-2 instruments, available at <http://www.iup.physik.uni-bremen.de/gome/gomemgii.html>.

The sunspot darkening index is obtained from data downloaded from NOAA NGDC, in the form of annual files of active region information compiled from the Air Force Solar Observing Optical Network (SOON) sites. The link is <http://www.ngdc.noaa.gov/stp/spaceweather.html>, ftp access, Solar\_Data, Sunspot\_Regions, USAF\_MWL. The files provide information about the areas, AS, and locations, μ, of individual sunspot active regions present on the solar disk.

Adding the total and spectral irradiance increments due to faculae and sunspots to the baseline reference total solar irradiance and spectral irradiance then determines the total and spectral irradiance at time, t. Values of the (invariant) total and spectral irradiance of the quiet Sun, and , are specified in the algorithm code as baseline references for this purpose. Solar spectral irradiance is estimated on a uniform wavelength grid, on 0.5 nm centers, in 1 nm bins from 115 .5 to 99999.5 nm, which is also the wavelength grid of the reference spectrum. The spectral irradiance is then summed into 3785 wavelength bins of variable width, designed appropriately for input to general circulation climate models.

Output files are generated separately for total solar irradiance and solar spectral irradiance. The TSI is one value per input facular brightening and sunspot darkening index values, typically daily. The units are W m-2. The SSI is one spectrum per input facular brightening and sunspot darkening index values, also typically daily. The solar spectral irradiance is produced in 3785 wavelengths binds, in units of W m-2 for the wavelength bin.

**Figure 1 Flow diagram of the Algorithm Processing.**



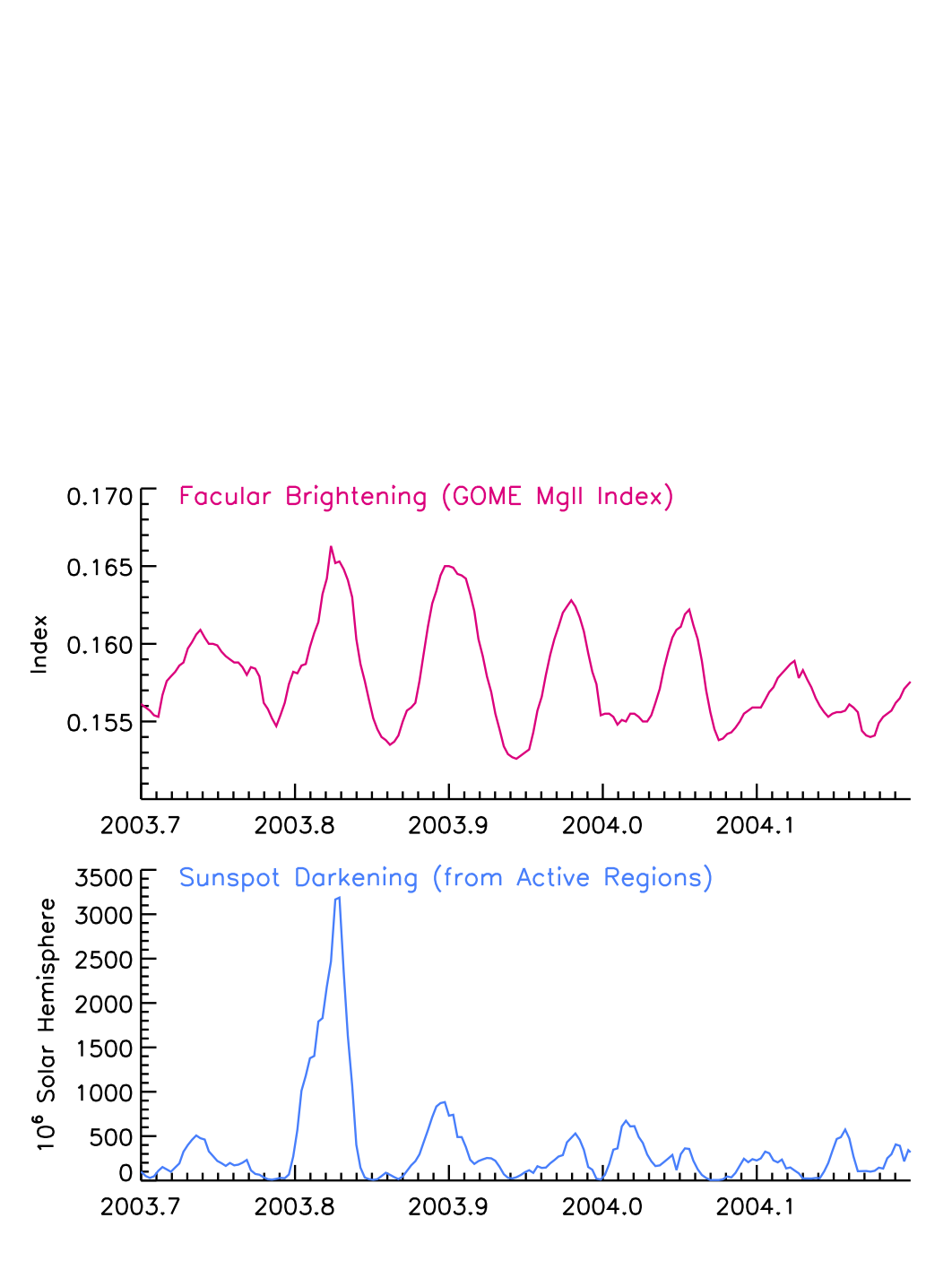
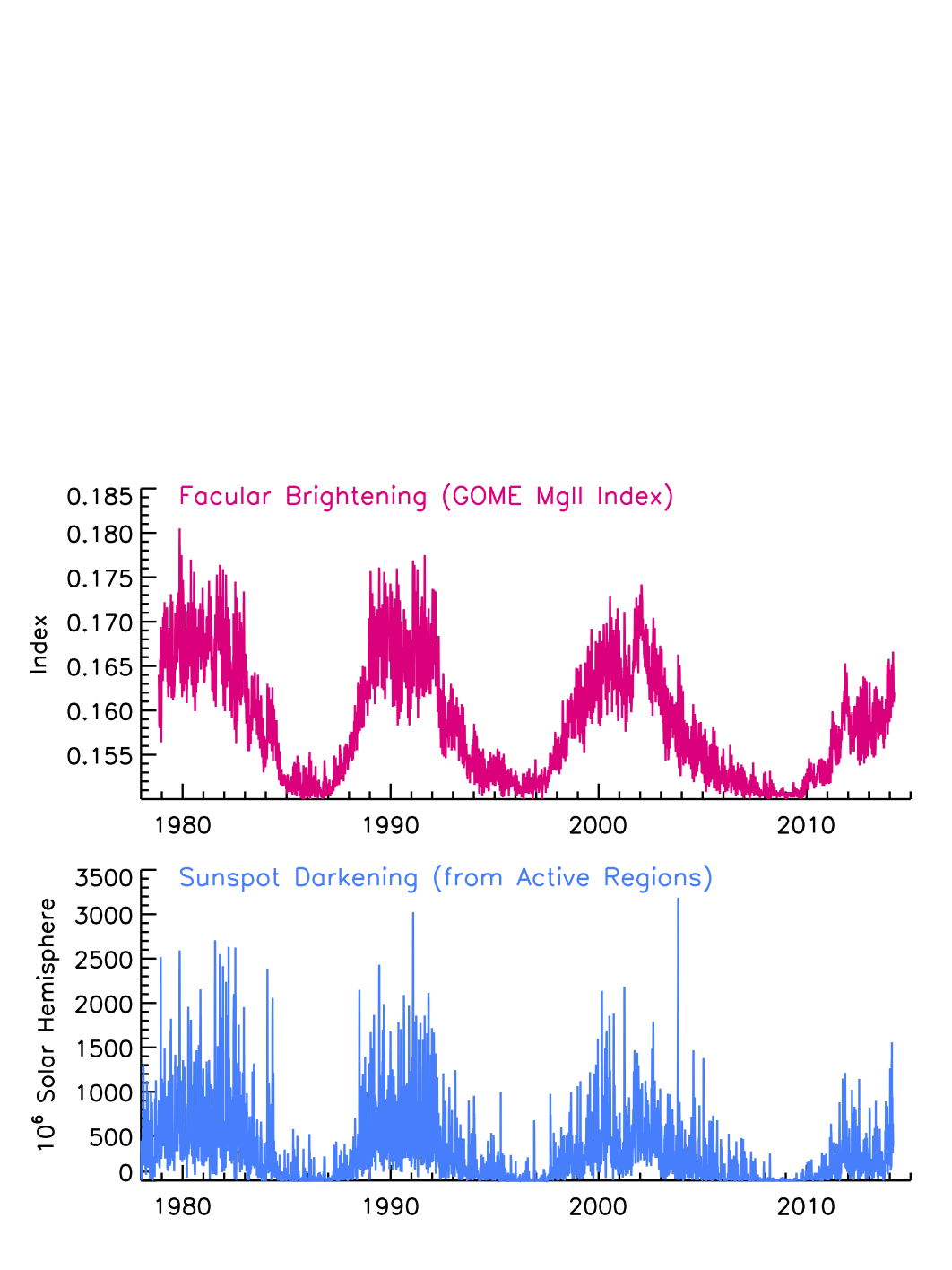
* 1. Algorithm Input
     1. Primary Input Data

Adopted for the facular brightening index, F(t), is the global (disk-integrated) Mg II emission from the Sun’s chromosphere, whose variations are dominated by chromospheric extensions of photospheric faculae. The Mg II index (which is the ratio of core emission in Fraunhofer lines to emission in the nearby continuum) is such a proxy because the core emission is enhanced in magnetically active bright regions, and the indices are sensitive indicators of the total (net) emission from all bright regions on the solar disk (Skupin et al., 2004; Snow et al., 2005). Furthermore, as the ratio of absolute fluxes, the Mg II index is (in principle) less susceptible to instrumental sensitivity changes that potentially contaminate the temporal fidelity of time series.

The sunspot darkening index, S(t), is calculated following Lean et al. (1998) using direct observations of the areas, AS and heliocentric locations, μ, of Nspot individual sunspot regions that the Air Force SOON (and other ground- and space-based) sites observe daily. The calculation effectively sums the projected area of sunspot regions on the solar hemisphere and multiplies this by the contrast of sunspots relative to the background (reference) Sun, taking into account variations with limb position on the solar disk. Areas prior to 1976 are reduced by 20% to account for systematic area differences between Greenwhich and Air Force SOON measurements.

Summing over all sunspots, and including an empirically-determined dependence of sunspot contrast on area (Brandt et al., 1994) the sunspot darkening index is

where μ = cos(latitute) × cos(longitude) for spot latitude (adjusted for the Bo angle of the Sun’s axis to the ecliptic plane) and longitude in heliocentric coordinates.

Figure 2 shows time series of the facular brightening and sunspot darkening indices since 1978, illustrating their evolution during the solar cycle as active regions emerge, evolve, and disappear from the solar disk. Figure 3 shows variations of these indices during 2003-2014, illustrating their shorter-term modulation arising from solar rotation. The facular brightening, F(t), and sunspot darkening, S(t), time-varying inputs that the Solar Irradiance Climate Data Record algorithm uses to calculate solar irradiance must be carefully calculated and validated, with possible erroneous values flagged for investigation.

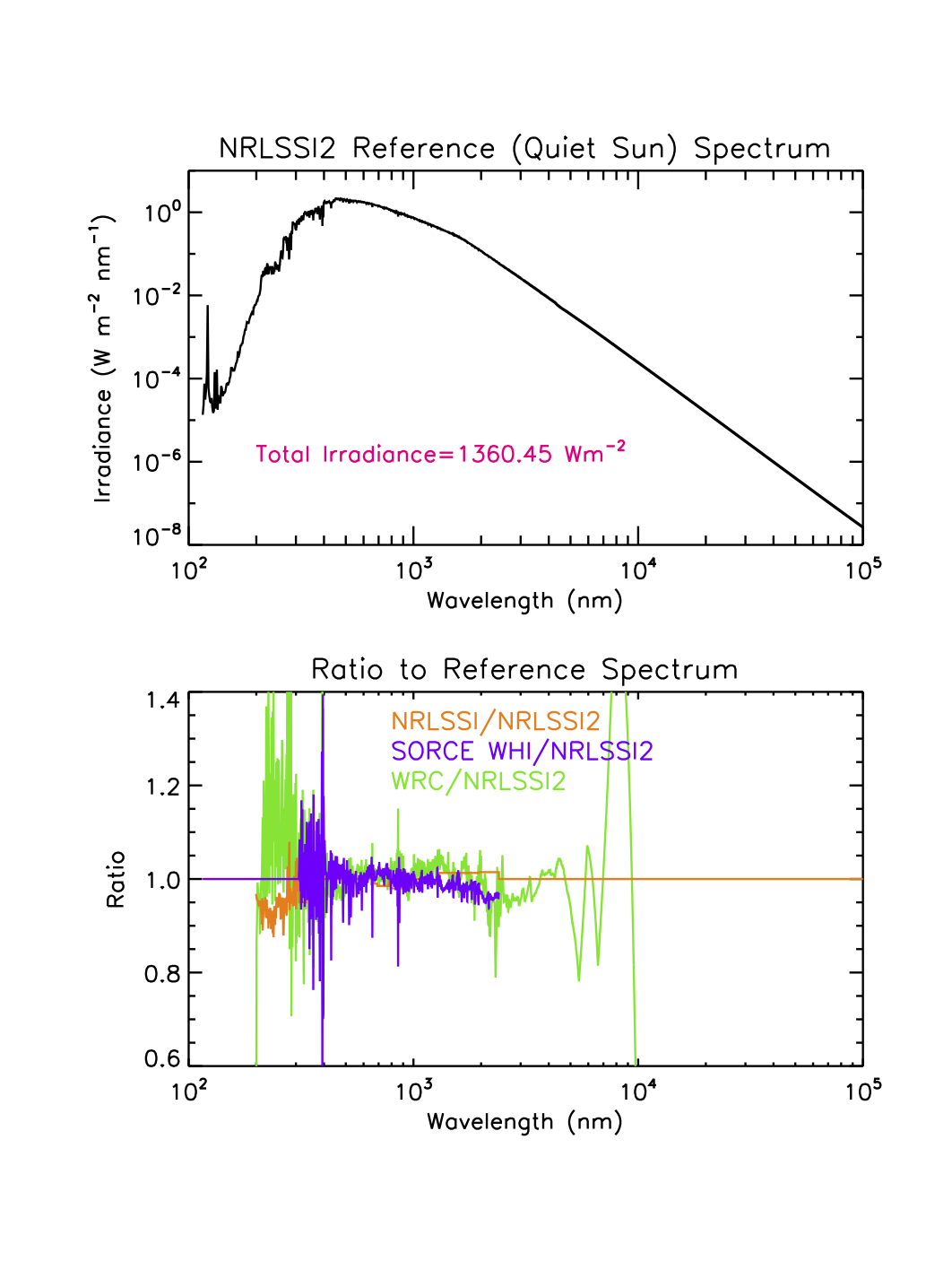
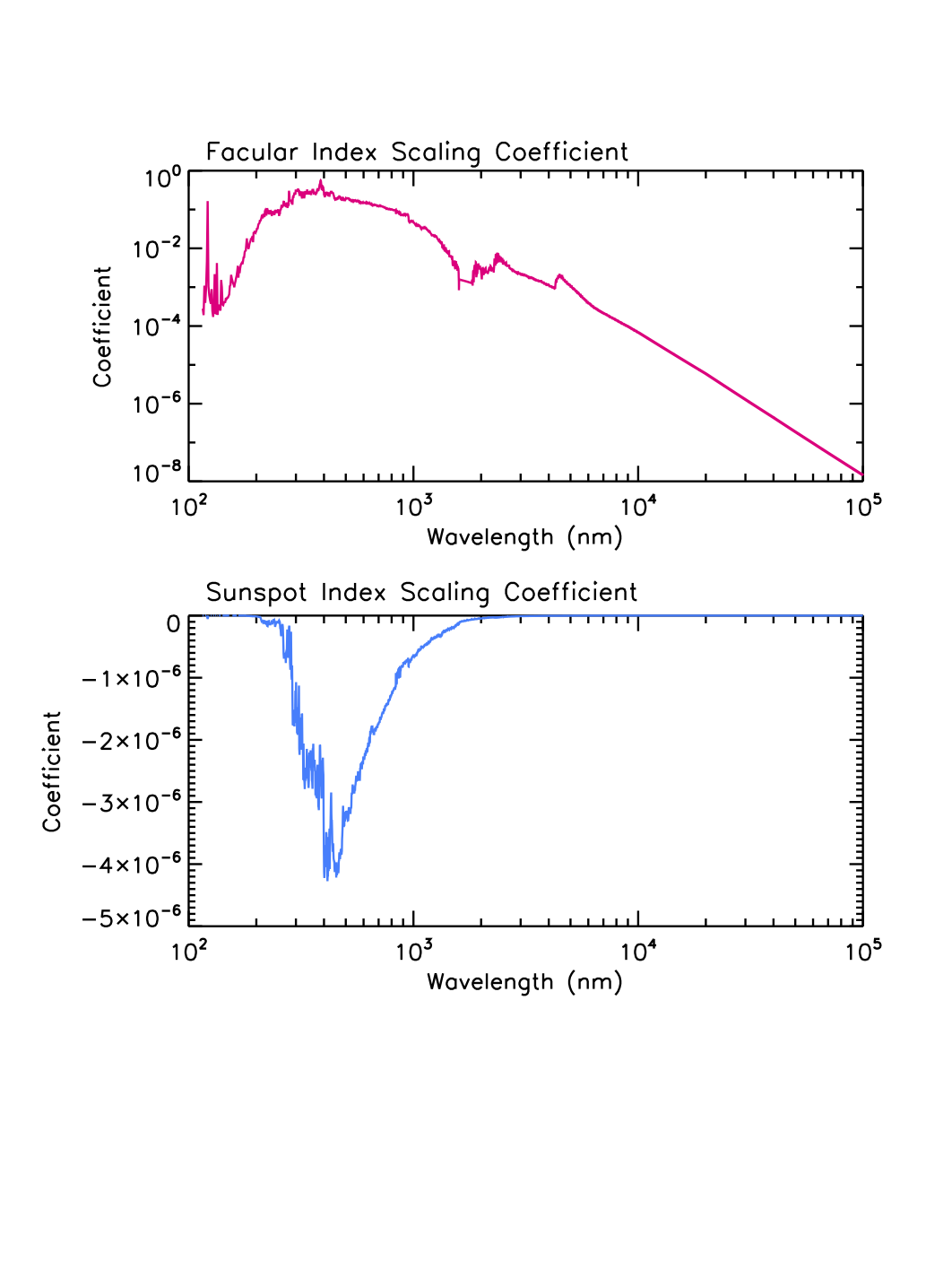
**Figure 3 Facular brightening and sunspot darkening time series during solar rotation.**

Figure 2 Facular brightening and sunspot darkening time series during the solar cycle.

* + 1. Ancillary Data

The algorithm uses constant baseline values of the total solar irradiance, TQ, and solar spectral irradiance, IQ (λ), to represent the quiet sun, as shown in Figure 4. Constant bolometric and wavelength-dependent coefficients, shown in Figure 5, linearly scale the facular brightening and sunspot darkening inputs to produce corresponding irradiance increments that are then applied to adjust the baseline irradiance, either increasing or deceasing it depending on the wavelength-dependent strengths of the facular and sunspot influences at that time. The bolometric and wavelength-dependent coefficients are included as part of the algorithm. Additional (invariant) data also specified in the algorithm are the angles of the solar rotation axis and the ecliptic plane, termed the beta angle, throughout the year. This angle is used to adjust the projection of sunspot areas to the direction of the Earth, as it orbits the Sun.

New versions of the model are generated when and if the coefficients are determined to need revision, based on additional observations and analysis sufficient to permit reformation of the model coefficients.



**Figure 5 Scaling coefficients that convert the facular brightening and sunspot darkening indices in Figures 2 and 3 to their equivalent irradiance change, in energy units.**

**Figure 4 Reference Spectrum of the quiet Sun in NRLSSI2**

* + 1. Derived Data

Once the facular brightening and sunspot blocking inputs are obtained, they are linearly scaled to convert them to equivalent increments of total and spectral irradiance, which are then added to the baseline (quiet) total and spectral irradiance to determine solar irradiance at time t as specified in 3.2. The application of the irradiance increments to the baseline spectral irradiance is implemented in 1 nm bins on 0.5 nm grid centers. The 1 nm spectral irradiance thus calculated is then summed into 3785 wavelength bins of varying width:

1 nm bins on wavelengths grid centers from 115.5 to 749.5 nm (635 bins)

5 nm bins on wavelength grid centers from 752.5 to 4997.5 nm (850 bins)

10 nm bins on wavelength grid centers from 5005.0 to 9995.0 nm (500 bins) 50 nm bins on wavelength grid centers from 10025.0 to 99975.0 nm (1800 bins)

There are no further processing steps; the calculated total and spectral irradiance are then written to output files. A time series of the output total solar irradiance since 1978 is shown in Figure 6 and examples of the corresponding solar spectral irradiance outputs in broad wavelength bands are shown in Figure 7.

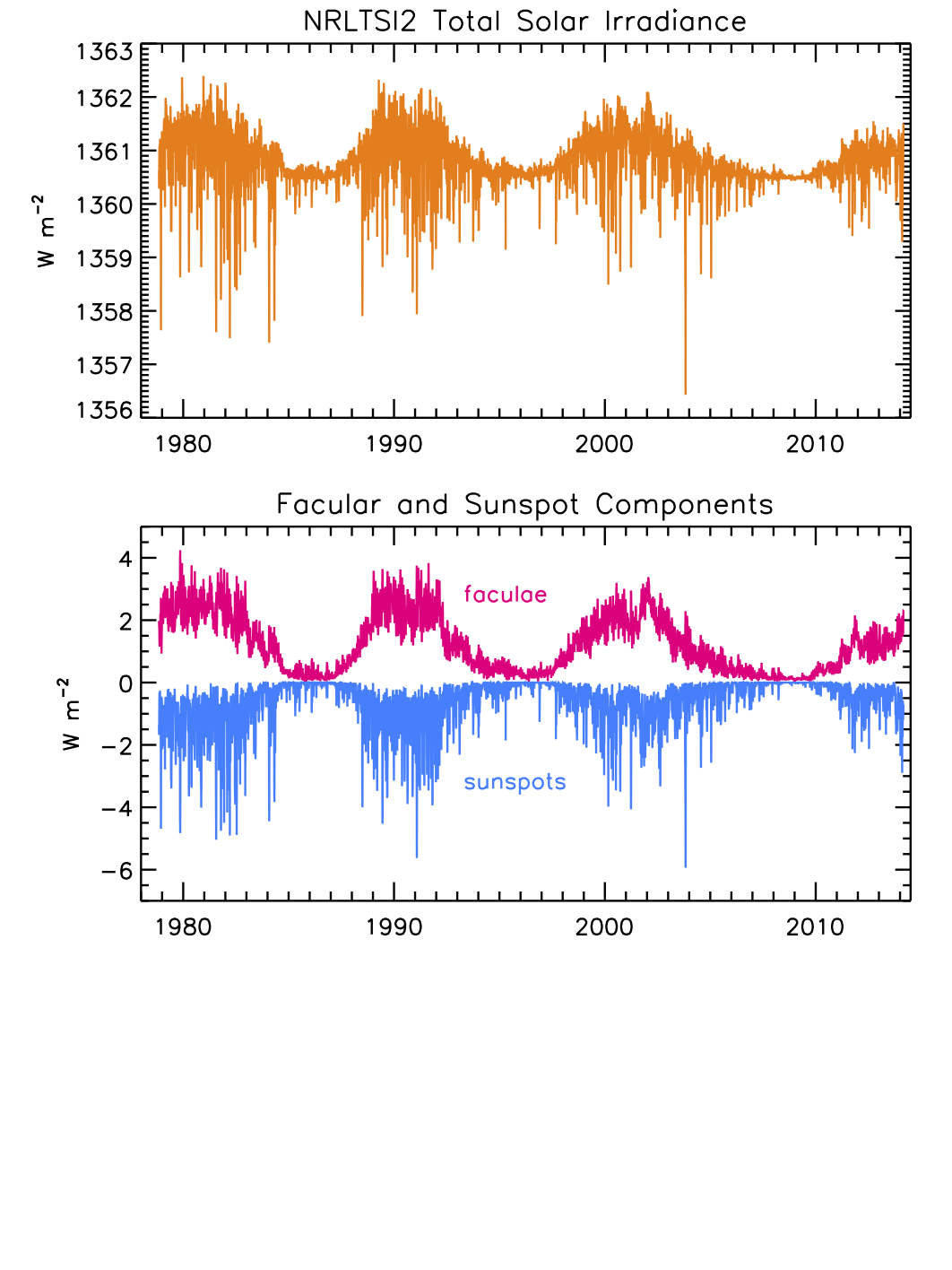
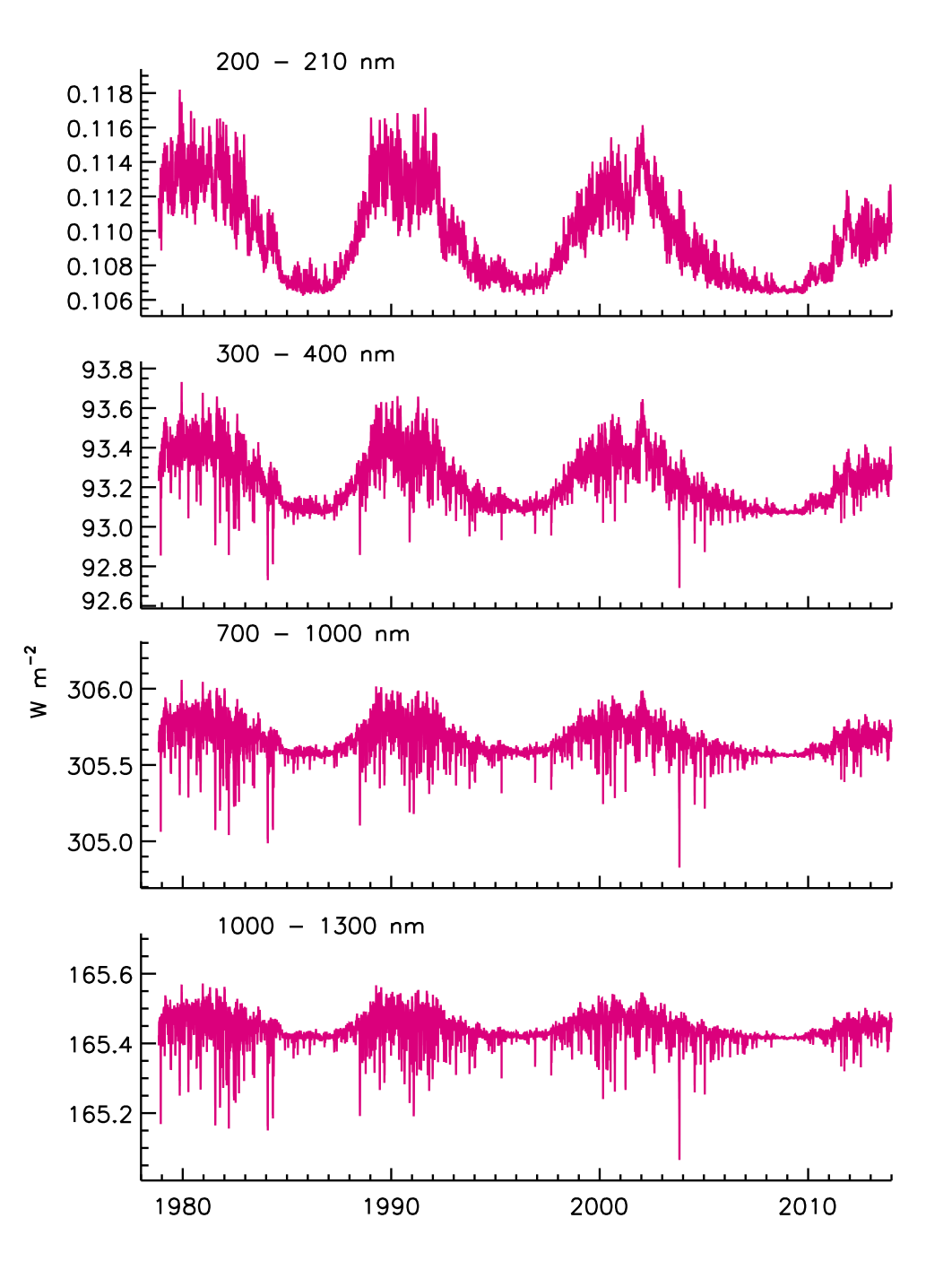
Because validation of the facular brightening and sunspot darkening inputs is crucial for reliable solar irradiance calculations, the processed facular brightening and sunspot darkening time series (e.g., Figure 2) used in the algorithm are output to a separate file, to facilitate validation with independent datasets.

Figure 7 Spectral irradiance in broad wavelength bands reconstructed with the NRLSSI2 model.

Figure 6 Total solar irradiance reconstructed with the NRLTSI2 model (upper), and the corresponding the facular and sunspot contributions (lower).

* + 1. Forward Models

Not applicable.

* 1. Theoretical Description

The overall approach of the Solar Irradiance Climate Data record algorithm builds on, and advances, the NRLTSI and NRLSSI models, as described in Lean (2000), Lean et al. (2005) and summarized in Lean and Woods (2010). Developed over a decade ago, prior to the launch of the Solar Radiation and Climate Experiment (SORCE) spacecraft (Rottman et al., 2005), the NRLTSI and NRLSSI solar irradiance variability models have been widely used for a variety of model simulations of climate and atmospheric change, including for the IPCC reports (e.g., Schmidt et al., 2011), and compared as well with other solar spectral irradiance variability models (Thuiller et al., 2013).

**Original NRLTSI and NRLSSI Models**

The NRLTSI model was formulated using a composite of total solar irradiance constructed by Fröhlich and Lean (2004), by combining observations made by Nimbus 7, ACRIM on SMM and UARS and PMOD on SOHO. The NRLSSI model was constructed for wavelengths less than 400 nm from a linear association of spectral irradiance variations observed by the Solar Stellar Irradiance Comparison Experiment (SOLSTICE, Rottman, 2000), relative to a reference spectrum (the average SOLSTICE spectrum during the UARS time period), with corresponding changes in facular brightening and sunspot darkening, also relative to their respective reference values.

Lacking observations of solar spectral irradiance variability at wavelengths longer than ~400 nm, NRLSSI’s spectral irradiance variations in the visible and infrared spectral regions were determined from the wavelength-dependence of the sunspot and facular contributions, according to their respective theoretical contrasts (ratio of emission to the background quiet solar atmosphere) determined in a theoretical solar atmosphere model (Unruh et al., 2000). For the quiet irradiance spectrum a composite was compiled on a 1 nm grid from space-based observations made by SOLSTICE on UARS (from 120 to 401 nm) and SOLSPEC on the ATLAS shuttle mission (from 0401 to 874 nm) (Thuillier et al., 1998), and a theoretical spectrum at longer wavelengths (Kurucz, 1991). The agreement among these three spectra in their regions of overlap is better than 2%, which is well within their absolute measurements uncertainties (Thuillier et al., 1998). The initially compiled composite spectrum was multiplied by 0.99 at all wavelengths to make its integral equal the independently measured total irradiance of the quiet Sun, whose most likely value at that time was considered to be 1365.5 Wm-2.

Neither Lean (2000) nor Lean et al (2005) refer explicitly to the spectral irradiance calculations, described above, as the NRLSSI model. Following the extension of the model to include the extreme ultraviolet spectrum (Lean et al., 2011), the designation NRLSSI was chosen (e.g., as summarized in Lean and Woods, 2010) to collectively describe an empirical capability to specify the entire solar spectral irradiance and its variability from 1 to 100,000 nm. The NRLSSI model calculates the solar spectral irradiance in 1 nm bins across the entire electromagnetic spectrum, daily since 1950, monthly since 1882 and annually since 1610.

**Newly Formulated NRLTSI2 and NRLSSI2 Models**

The Solar Irradiance Monitor (SIM) and the SOLSTICE instruments on SORCE have provided new observations of spectral irradiance variability throughout the descending phase of solar cycle 23 and the ascending phase of cycle 24. New versions of irradiance variability models, designated NRLTSI2 and NRLSSI2, have been formulated directly from the TIM and SIM observations, and are implemented in the Solar Irradiance Data record algorithm that this C-ATBD describes. The measurements made by the Total Irradiance Monitor (TIM) on SORCE indicate that the actual total irradiance of the quiet sun is ~5 Wm-2 lower than in NRLTSI (Kopp and Lean, 2011), and the NRLTSI2 and NRLSSI2 models are consistent with this new, lower value of solar irradiance.

As described in Section 3.2, the basic formulation of total and spectral solar irradiance, T(t) and I(λ,t), determines their variations arising from faculae and sunspots superimposed on specified, invariant, quiet sun reference values, TQ(t) and IQ(λ,t) (Figure 4), as,

where and are functions of the facular brightening index, F(t), and and are functions of the sunspot darkening index, S(t). Specifically, for the NRLTSI2 total irradiance variability model

where FQ and SQ (= 0) are the values of the facular brightening and sunspot darkening indices corresponding to TQ, i.e., for the quiet sun.

Similarly, the corresponding facular brightening and sunspot darkening components in the NRLSSI2 spectral irradiance variability model are

The a, b, c(λ) and d(λ) coefficients for faculae and sunspots are specified (Figure 5) and supplied with the algorithm. Note that the a and c coefficients are nominally zero so that when F=FQ and S=SQ then T=TQ and I=IQ. The additional wavelength-independent terms in the spectral irradiance facular and sunspot components evaluated with the eF and eS coefficients provide small adjustments to ensure that numerically

The F(t) and S(t) are calculated using independent solar observations made approximately daily, respectively, the Mg index of global - disk integrated - facular emission and information about the area and locations of sunspot active regions on the solar disk, as describe in 3.3.1. Ideally, were the F(t) and S(t) physically and observationally “perfect” indicators of the sunspot and faculae sources at each wavelength, then the coefficients eF and eS would be zero. Improvements in F(t) and S(t) may enable this in future versions of the algorithm.

The NRLTSI2 model uses multiple linear regression to determine the scaling coefficients of the facular brightening and sunspot darkening time series that best reproduce the total solar irradiance variability measure directly by TIM from 2003 to 2014:

The observed, TTIMl and modeled, Tmod,total solar irradiance have a correlation coefficient of 0.96 and the standard deviation of the residuals, TTIM – Tmod, is 0.1 Wm-2.

Because SIM’s calibration is far less stable than TIM’s, it is likely that instrumental trends are present in SIM’s solar spectral irradiance measurements (Lean and DeLand, 2012). This precludes the formulation of reliable models of solar spectral irradiance from the SIM observations, directly. Instead, a relationship of solar spectral irradiance variability to sunspot darkening and facular brightening is first determined using observations of solar rotational modulation: instrumental trends are smaller over the (much) shorter rotational times scales than during the solar cycle. For each 1 nm bin, the observed spectral irradiance and the facular brightening and sunspot darkening indices are detrended by subtracting 81-day running means. Multiple linear regression is then used to determine the relationships of the detrended time series:

The range of facular variability in the detrended time series is smaller than during the solar cycle which, together with the “imperfect” nature of the facular brightening (and sunspot darkening) index, causes the coefficients of models developed from detrended time series to differ from those developed from direct (i.e., not detrended) observations.

The total irradiance observations are used to numerically determine ratios of the coefficients obtained from multiple regression using direct observations of solar irradiance, with those obtained from multiple regression of detrended observations. A second model of the TIM observations was formulated analogous to that used for NRLTSI2, but using detrended, instead of direct, time series. The ratios of the coefficients for the two different approaches were then used to adjust the coefficients for spectral irradiance variations, determined from the detrended time series. For wavelengths longer than 295 nm, where both sunspots and faculae modulate the spectral irradiance (as they do total irradiance), dS and dF are estimated as

For wavelengths shorter than 295 nm where faculae are the dominant cause of irradiance variability (and dS(λ) ~ 0), the adjustments for the coefficients were estimated using the Ca K time series, a facular index that is independent of the Mg II index, and a well-recognized indicator of UV spectral irradiance variability.

The average difference between the spectrally integrated NRLSSI2 model and NRLTSI2 from 2003 to 2014, is 0.016 Wm-2 and the standard deviation of the differences is 0.012 Wm-2.

**Speculated Irradiance Changes beyond the Solar Cycle**

As well as being a dominant determinant of solar cycle irradiance variations, faculae are speculated to cause longer-term (decadal to centennial) irradiance trends. Specifying the past evolution of the facular signal is therefore necessary for reconstructing historical irradiance variations. But unlike the sunspot signal, which is suggested by direct observations of sunspot numbers, the facular component is highly uncertain and dependent on circumstantial evidence. For example, based on current observations of facular contrast and disk coverage, the disappearance of all faculae from the Sun's surface is estimated to decrease total solar irradiance about 0.1% (Lean et al., 1992). Attempts have been made to translate variations in the chromospheric activity in Sun-like stars to a plausible range of the facular influence on solar irradiance (Lean et al., 1992, 1995), with results broadly consistent with inferences from the cosmogenic and geomagnetic indices. Changes in solar structure are also considered as possible sources of long-term irradiance variations in addition to, or instead of, facular variations (Hoyt and Schatten, 1993; Tapping et al., 2007) producing levels as much as 0.3% below contemporary solar minima values (e.g., review of Maunder Minimum levels in Lean et al., 2005).

In the original version of NRLSSI (Lean, 2000) the long-term “background” component of the facular index, FBG(t), was specified as a 15-year running mean of annual sunspot group numbers in which the reduction from the quiet Sun to the Maunder Minimum is 92% of the increase in FBG(t) from the quiet Sun to cycle maximum (Nov. 1989). These changes mimicked the reduced Ca fluxes in non-cycling Sun-like stars compared with the range of fluxes in cycling Sun-like stars (Radick et al, 1998 and Lean et al., 2001, provide additional details) which at the time of the NRLSSI model were thought to exemplify long-term solar irradiance changes. However, a subsequent reassessment of the stellar data was been unable to recover the original bimodal separation of (lower) Ca emission in non-cycling stars (assumed to be in Maunder Minimum type states) compared with (higher) emission in cycling stars (Hall and Lockwood, 2004). Nor do long-term trends in the aa index and cosmogenic isotopes (generated by open flux) necessarily imply equivalent long-term trends in solar irradiance (which track closed flux) according to simulations of the transport of magnetic flux on the Sun and propagation of open flux into the heliosphere (Lean et al., 2002; Wang et al., 2005).

These developments motivated revision of the long-term “background” component of the NRLSSI model using a flux transport model to estimate the plausible magnitude of a long-term secular facular component. The flux transport model (with variable meridional flow) simulates the eruption, transport, and accumulation of magnetic flux on the Sun’s surface from the Maunder Minimum to the present in strengths and numbers proportional to the sunspot number (Wang et al., 2005). The model estimated variations in both open and total flux arising from the deposition of bipolar magnetic regions (active regions) and smaller-scale ephemeral regions on the Sun’s surface: The open flux compares reasonably well with the geomagnetic and cosmogenic isotopes, which gives confidence that the approach is plausible. A small accumulation of total flux (and possibly ephemeral regions) produces a net increase in facular brightness that, in combination with sunspot blocking, permits the reconstruction of total solar irradiance.

The increase in total solar irradiance from the Maunder Minimum to the present-day quiet Sun is about 0.04%, based on the flux transport model simulations. For comparison, because of the larger background facular component adopted in Lean (2000), the increase from the Maunder Minimum to the present-day quiet Sun in the original version of the NRLSSI model is about 0.16%, four times larger. (See Lean et al., 2005, for comparison of different estimates of TSI from the Maunder minimum to the present). The spectral irradiance changes in the NRLSSI2 model are consistent with the Wang et al. (2005) flux transport simulations and were obtained by using a background component 27% of that adopted in the spectral irradiance reconstructions of Lean (2000).

* + 1. Physical and Mathematical Description

The NRLTSI2 and NRLSSI2 models, which the Solar Irradiance Climate Data record algorithm utilizes, assume that bright faculae and dark sunspots are the only causes of solar irradiance variability on contemporary time scales. The occurrence of these features on the Sun varies during the Sun’s 11-year activity cycle, producing a prominent 11-year cycle in solar irradiance. The rotation of the Sun on its axis alters the population of faculae and sunspots projected to earth, producing an additional 27-day irradiance modulation.

Following the approach of Lean et al (1998), when the sun is inactive, the “quiet” irradiance at the Earth (at a distance of 215 times the solar radius) is determined by integrating the radiance, R(λ,1), at the center of the disk (µ = 1) over the entire disk using the center-to-limb function, L(λ, µ) to define the ratio of R(λ, µ) to R(λ, 1) at heliocentric position µ, which ranges from 0 (at the disk’s limb) to 1 (at disk center). Thus the irradiance of the quiet sun is

Magnetic features – dark sunspots and bright faculae – when present on the solar disk solar disk alter the otherwise homogeneous distribution of radiance, and hence the irradiance at the Earth, which is at some (non-quiet) time, t, given as

where

is the ratio of the Sun’s radiance at heliocentric location, μ, relative to the radiance of the surrounding quiet Sun. This ratio is termed the contrast.

Separating radiance elements on the solar disk into those that are brighter than, darker than or equal to the quiet sun radiance permits expression of the irradiance as

where CF(λ,μ) and CS(λ,μ)are the contrasts of the faculae and sunspots, respectively. For the number of radiance elements defined as faculae, Nfac, and sunspots, Nspot , with actual area on the solar surface of Afac and Aspot, at a given time, t, the corresponding solar irradiance is (with A = dμ for solar radius rsun)

which is analogous *to* the basic formulation used in the Solar Irradiance Climate Data Record algorithm for the solar spectral irradiance:

where

The corresponding total solar irradiance is

which is analogous *to* the basic formulation used in the Solar Irradiance Climate Data Record algorithm for the total spectral irradiance:

The calculation of the sunspot darkening index in Section 3.3.1 is physically an estimate of made using the above theoretical basis with a number of assumptions and parameterizations, as follows. The center-to-limb variation is assumed to be independent of wavelength and specified as

The sunspot contrast CS(λ) = RS(λ) /RQ(λ) is assumed to be independent of μ and its bolometric (i.e., spectrally integrated) value, nominally 0.32, determined experimentally to depend on AS, where AS is expressed in millionths of the solar hemisphere, as reported by NOAA

With these assumptions,

Typically, information about sunspot areas and locations are recorded at different times throughout the day (depending on local time) by a dozen or so different ground-based stations. The sunspot darkening index used to evaluate NRLTSI2 and NRLSSI2, shown in Figure 6, is the average of all the available information on a given day. As well, individual sites calculate sunspot darkening factors, and it has also been applied to space-based white light images made by MDI on SOHO (F. Watson, http://www.nso.edu/staff/fwatson/STARA).

Although the facular brightening can be calculated similarly to the sunspot darkening index, as Lean et al. (1998) demonstrate for the irradiance at 200 nm using histograms of calibrated Ca K solar images to identify bright faculae, the characteristics of facular are in general poorly observed in solar imagery and inadequately specified compared with the more compact, darker and relatively well-defined sunspot regions. Furthermore, whereas sunspot regions are typically discrete and therefore relatively easily quantified, faculae occur with a continuous distribution of sizes and contrasts, so that statistical definitions (which can be ambiguous) are needed for practical quantification. Because of the lack of reliable quantitative data for facular areas, center-to-limb functions and contrasts, the NRLSSI2 model (like NRLSSI) calculates spectral irradiance change due to faculae as a linear function of a “flux” (i.e., disk-integrated) proxy of facular brightening, F(t).

* + 1. Data Merging Strategy

No data merging is needed for the NRLTSI2 and NRLSSI2 models to calculate solar irradiance.

* + 1. Numerical Strategy

The NRLTSI2 and NRLSSI2 models do not include numerical algorithms.

* + 1. Calculations

The algorithm calculates the total and solar spectral irradiance using an IDL procedure to applying the previously derived (and constant in time) coefficients to scale the two inputs, the facular brightening and sunspot darkening indices.

* + 1. Look-Up Table Description

There are no Look-Up tables.

* + 1. Parameterization

There are no parameterizations.

* + 1. Algorithm Output

**Total Solar Irradiance (TSI)**

The NRLTSI2 model produces a value of total solar irradiance on an absolute scale defined by TIM on SORCE, for given inputs of the facular brightening and sunspot darkening indices, estimated daily using inputs from ground and space-based solar observations when available, as specified above.

Typical total solar irradiance output files in NetCDF4 format have the structure as identified in Table 3. The files follow CF-1.5 metadata convections for variable names and attributes.

Table 3: Structure of NRLTSI2 output for daily values of TSI from 1978 to the present.

|  |  |  |  |  |
| --- | --- | --- | --- | --- |
| **Variable Name** | **Long Name** | **Standard Name** | **Units** | **Missing Value** |
| TSI | NOAA Fundamental Climate Data Record of Daily Total Solar Irradiance (Watt/ m\*\*2) | toa\_incoming\_shortwave\_flux | W m-2 | -99.0 |
| time | ISO-8601 date/time (YYYY-MM-DD) format | time | N/A | N/A |

The TSI data are organized in several different ways to support the projected needs of the user communities. Firstly, there is a 1-month file containing daily TSI, with approximately 1-3 month updates. Secondly, the daily TSI values are also be available in 1-year files. Thirdly, monthly-averaged values of TSI are available in 10-year files. Lastly, historical, annually-averaged TSI values are available in a single file.

The time ranges covered by the science data product is 1978 to the present for contemporary TSI, from 1882 to the present for historical, daily values and monthly-averaged TSI, and 1615 to the present for historical and annually-averaged TSI.

For the monthly-averaged and yearly-averaged files, the “Time” variable is year and month (yyyy mm), and year (yyyy), respectively.

The file naming conventions for the TSI data described above is as follows:

[data-short-name]\_[data-version]\_[version-revision]\_[extra-attribute]\_[begin-date]\_[end-date]\_[creation-date].nc

where [data-short-name] is ‘TSI’, [data-version] is the version number beginning at V001 indicating the initial release of the science data product, [version-revision] is the revision number beginning at R00 for revision 0, and [extra-attribute] is ‘day’ for daily TSI, ‘mon’ for monthly-averaged TSI, or ‘ann’ for annually averaged TSI. The [begin-date] and [end-date] denote the starting and ending time period for the science data product, while creation date denotes the date the file was created using the LASP LaTiS server (Section 5.1).

The [begin-date] and [end-date] formats vary based on whether the science data is daily TSI, or monthly or annually averaged. For example, the time format for daily TSI follows YYYYMMDD convention, monthly-averaged TSI follows YYYYMM convention, and annually-averaged TSI follows YYYY convention. For all files, the creation date follows a format of DDMMMYY.

**Solar Spectral Irradiance (SSI)**

The NRLSSI2 model produces solar spectral irradiance values on an absolute scale such that the integrated spectral irradiance is equivalent to the total irradiance observed by TIM on SORCE, for given inputs of the facular brightening and sunspot darkening indices, estimated daily using data from ground and space-based solar observations when available, as specified above.

Typical solar spectral irradiance output files in NetCDF4 format have the structure as identified in Table 4. The files follow CF-1.5 metadata convections for variable names and attributes.

Table 4: Structure of NRLSSI2 output for daily values of SSI from 1978 to the present.

|  |  |  |  |  |
| --- | --- | --- | --- | --- |
| **Variable Name** | **Long Name** | **Standard Name** | **Units** | **Missing Value** |
| SSI | NOAA Fundamental Climate Data Record of Daily Solar Spectral Irradiance (Watt m-2 nm-1) | N/A | W m-2 nm-1 | -99.0 |
| Central Wavelength | Wavelength grid center | N/A | nm | N/A |
| Wavelength Bands | Wavelength bands. Centered on Central Wavelength | N/A | nm | N/A |
| TSI | NOAA Fundamental Climate Data Record of Daily Total Solar Irradiance (Watt m-2) | toa\_incoming\_shortwave\_flux | W m-2 | -99.0 |
| time | ISO-8601 date/time (YYYY-MM-DD) format | time | N/A | N/A |

The SSI data are organized in different ways to support the projected needs of the user communities. Firstly, there is a 1-month file containing daily SSI, with approximately 1-3 month updates. Secondly, the daily SSI values is available in 1-year files. Thirdly, monthly-averaged values of SSI are available in 1-year files. Lastly, annually-averaged SSI values are available in 10-year files.

The time ranges covered by the science data product is 1978 to the present for contemporary SSI, from 1882 to the present for historical, daily values and monthly-averaged SSI, and 1615 to the present for historical, annually-averaged SSI.

For the monthly-averaged and yearly-averaged files, the “Time” variable is year and month (yyyy mm), and year (yyyy), respectively.

The file naming conventions for the SSI data described above is as follows:

[data-short-name]\_[data-version]\_[version-revision]\_[extra-attribute]\_[begin-date]\_[end-date]\_[creation-date].nc

where [data-short-name] is ‘SSI’, [data-version] is the version number beginning at V001 indicating the initial release of the science data product, [version-revision] is the revision number beginning at R00 for revision 0, and [extra-attribute] is ‘day’ for daily SSI, ‘mon’ for monthly-averaged SSI, or ‘ann’ for annually averaged SSI. The [begin-date] and [end-date] denote the starting and ending time period for the science data product, while creation date denotes the date the file was created using the LASP LaTiS server (Section 5.1).

The [begin-date] and [end-date] formats vary based on whether the science data is daily SSI, or monthly or annually averaged. For example, the time format for daily SSI follows YYYYMMDD convention, monthly-averaged SSI follows YYYYMM convention, and annually-averaged SSI follows YYYY convention. For all files, the creation date follows a format of DDMMMYY.

The science data product for solar spectral irradiance also contains the value of TSI because users interested in SSI (for example, climate modelers) also require the integrated quantity to constrain the total incoming solar irradiance. By including the TSI with the SSI science data product, we provide the user community the necessary data in a single file.

1. Test Datasets and Outputs
   1. Test Input Datasets

As a first test that the algorithm is performing as expected, the total solar irradiance is compared numerically with the integral of the solar spectral irradiance, and the individual faculae and sunspot components are also compared. Figure 8 shows these comparisons. For specific versions of the solar irradiance algorithm, the comparison of NRLTSI2 and ΣNRLSSI2 may have slightly different offsets and standard deviations, as a result of altered facular and sunspot, wavelength-dependent coefficients. The values for the mean and standard deviation of the residual time series NRLTSI2 and ΣNRLSSI2 given in Figure 8 corresponds to the initial version of the algorithm.

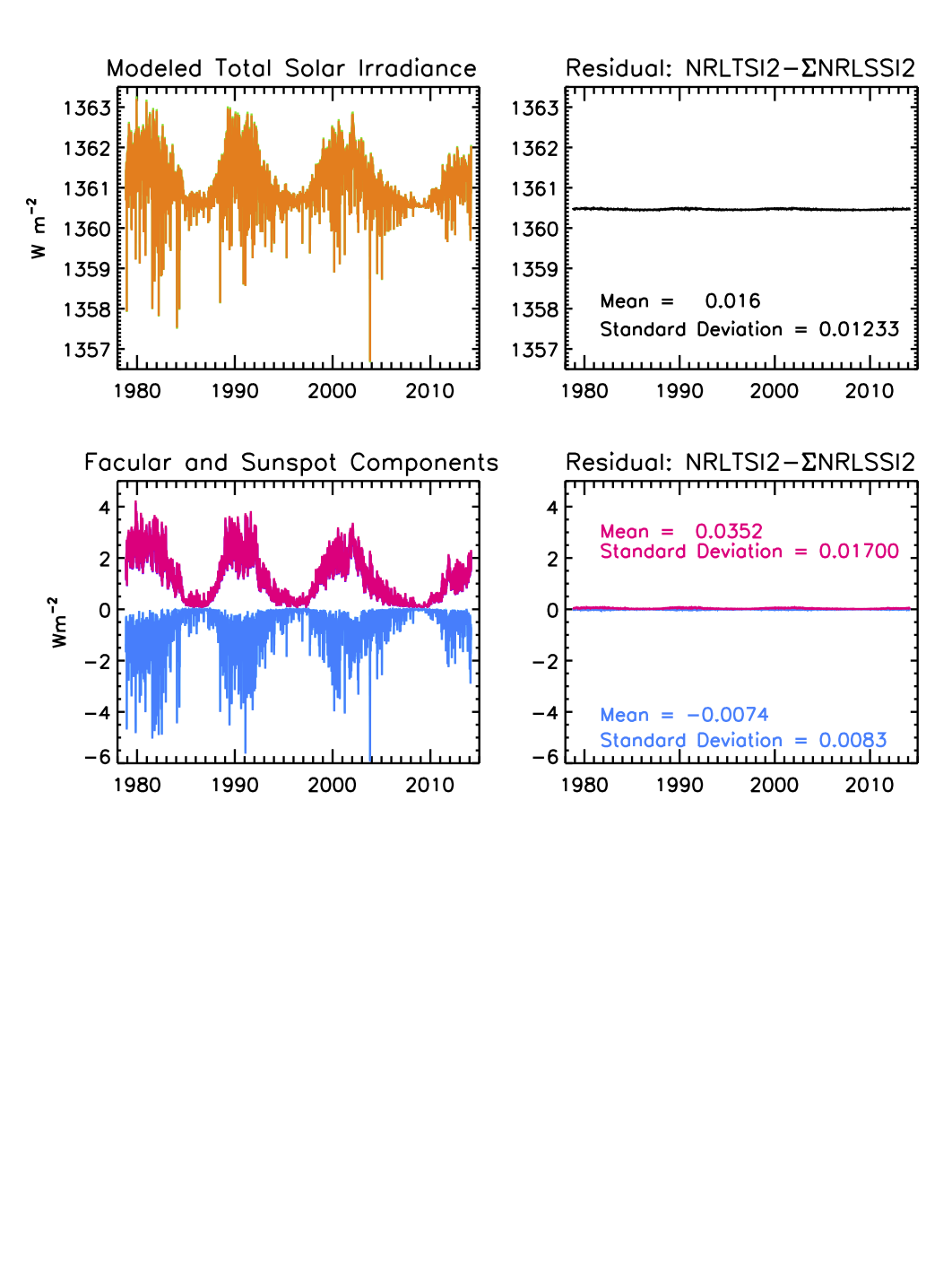
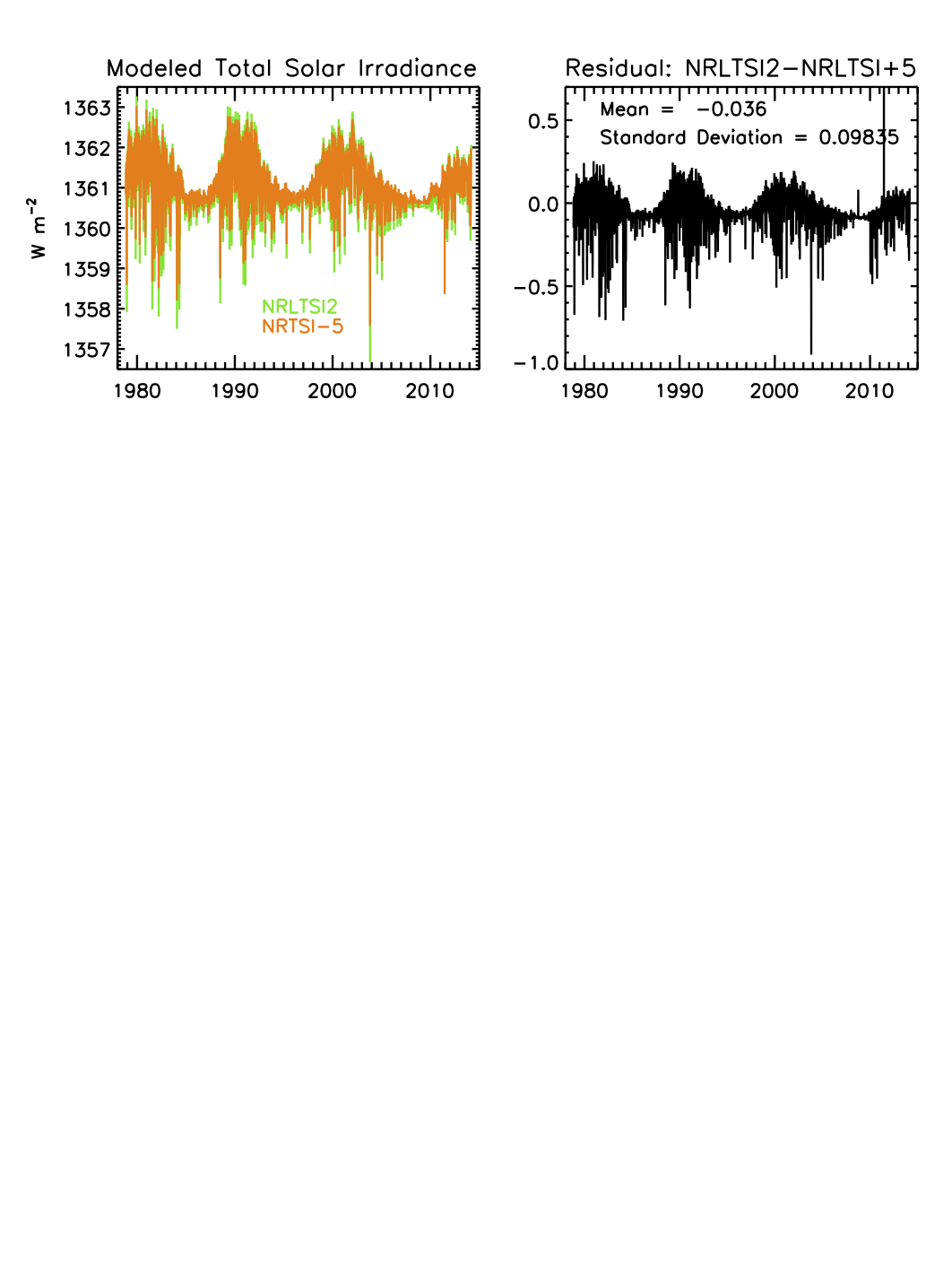
A second test for algorithm performance is comparison of selected time series of the NRLTSI2 and NRLSSI2 irradiance values that the algorithm calculates with the corresponding time series of the original NRLTSI and NRLSSI models (which have, themselves, been widely compared with available observations). Figure 9 shows such a comparison for the total solar irradiance. NRTSI2, constructed directly from the SORCE/TIM observations, has a lower absolute scale and slightly larger variability than that of NRLTSI. Figure 10 compares the spectral irradiance in selected broad wavelength bands according to NRLSSI2 with the corresponding values on NRLSSI.

Figure 9 Comparison of NRLTSI2 total solar irradiance, calculated by the C-ATBD algorithm with an earlier model, NRLTSI.

Figure 8 Comparison of NRLTSI2 total irradiance calculated by the C-ATBD algorithm, with the integral of the NRLSSI2 spectra also calculated by the algorithm, upper, and of their respective facular and sunspot components (lower).

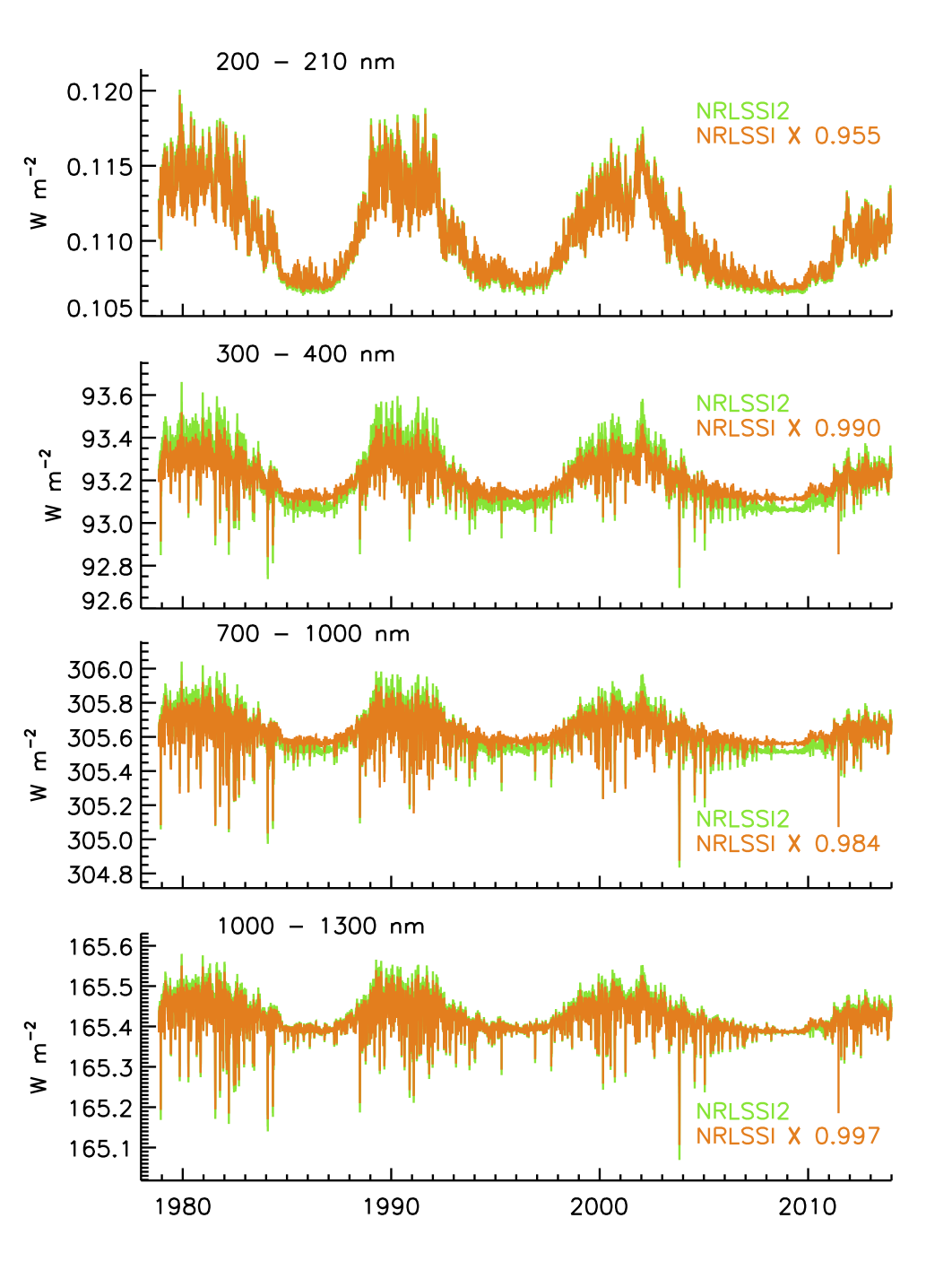
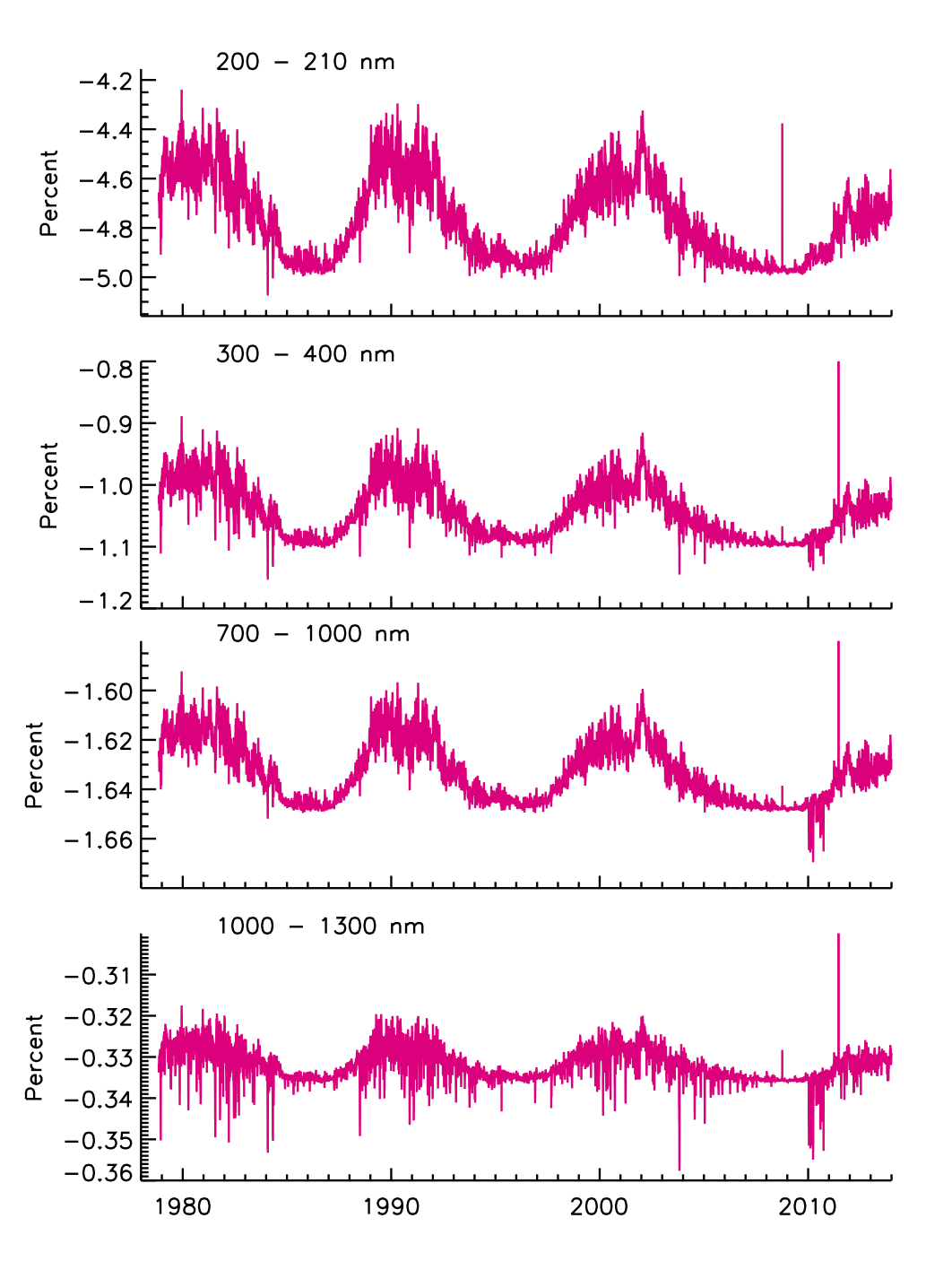


Figure 10: Comparison of NRLSSI2 solar spectral irradiance in the selected broad wavelength bands shown in Figure 7, calculated by the C-ATBD algorithm, with an earlier model, NRLSSI. On the left are the time series and on the right are the percentage differences.

* 1. Test Output Analysis
     1. Reproducibility

The algorithm’s calculation of total and spectral irradiance is 100% numerically reproducible given identical sunspot darkening and facular brightening inputs.

* + 1. Precision and Accuracy

Numerically, the algorithm itself precisely and accurately calculates the total and spectral irradiance according to the specified sunspot darkening and facular brightening inputs, the baseline (quiet) reference irradiance values and the bolometric and wavelength-dependent scaling factors. An immediate verification of the algorithm performance is the numerical comparison of the total solar irradiance, T(t), with the integral of the solar spectral irradiance, , which should agree to within a known amount, i.e., . Similarly, the separate facular and sunspot components that the algorithm calculates for the total irradiance and the integrated spectral irradiance should agree to within known amounts, i.e., and .

The precision and the accuracy of the derived solar irradiance depend on:

1) uncertainties in the absolute scale of the reference quiet sun values, which are obtained from direct solar irradiance measurements made by instruments whose calibration is traceable to NST standards (Kopp and Lean, 2011).

2) uncertainties in the input facular brightening and sunspot darkening values, including those related to measurements of the indices, and to assumptions about the indices’ representation of facular brightening and sunspot darkening (see Table 5 for a listing).

3) statistical uncertainties on the coefficients used to scale the facular and sunspot inputs to equivalent irradiance increments.

4) assumptions used to formulate the basic algorithm equations (see Table 5 for a listing).

* + 1. Error Budget

The original NRLTSI and NRLSSI values for modeled solar irradiance did not have accompanying error estimates. Table 5 provides Initial time-independent uncertainty estimates for the first three sources identified in 4.2.2. Uncertainties in the absolute scale of the irradiance are those reported for the direct measurements. Uncertainties associated with the solar indices inputs to the algorithm are obtained from statistical analysis of the index time series and their input data, and from comparisons with other similar indices. Uncertainties in the coefficients that transform the input indices to irradiance are obtained form the statistical output of the regression analysis used to construct the model that the algorithm uses.

Uncertainties arising from the assumptions used to formulate the algorithm are more difficult to assess objectively and establish quantitatively. Future work in support of ongoing efforts to produce a robust solar irradiance climate data record will extend the initial error estimates for the NRLTSI2 and NRLSSI2 modeled solar irradiance given in Table 5, including their time and wavelength dependencies. The future uncertainty estimates will incorporate an understanding of the impacts of the assumptions in the algorithm’s theoretical basis and uncertainties in the input facular brightening and sunspot darkening values on the derived solar irradiance (itemized in Table 6). This future understanding will reflect previous peer-reviewed studies and statistical results from the operational production of the modeled solar irradiance.

Table 5: Estimated uncertainties in the spectral irradiance (in percent) produced by the algorithm.

|  |  |  |  |  |  |
| --- | --- | --- | --- | --- | --- |
| **Source of Uncertainty** | **Total Irradiance** | **Spectral Irradiance**  **250 nm** | **Spectral Irradiance**  **500 nm** | **Spectral Irradiance**  **1000 nm** | **Spectral Irradiance**  **2000 nm** |
| absolute calibration scale | 0.03 | 5 | 3 | 3 | 5 |
| facular index |  |  |  |  |  |
| sunspot index |  |  |  |  |  |
| statistical uncertainties in model facular index coefficient |  |  |  |  |  |
| statistical uncertainties in model sunspot index coefficient |  |  |  |  |  |
| rotational to solar cycle scaling coefficients | N/A |  |  |  |  |
| combined uncertainty |  |  |  |  |  |

1. Practical Considerations
   1. Numerical Computation Considerations

The Solar Irradiance Data Record algorithm uses basic algebra. There are no matrix inversions, extrapolations, or interpolations in the algorithm itself, which is computationally rapid, efficient and repeatable.

The Solar Irradiance Data Record processing utilizes the LASP Time Series Server (LaTiS), an Application Programming Interface that allows software programs to access modeled solar irradiance data sets on LASP’s Interactive Solar Irradiance Datacenter (LISIRD) for a desired time range and for the desired variables. As well, the Solar Irradiance Data Record is provided to NOAA NCDC as NetCDF4 data files (Section 3.4.7).

The Solar Irradiance Data Record processing system accesses data files of sunspot area and location data of individual sunspot active regions present on the solar disk stored at NOAA’s National Geophysical Data Center (NGDC). The automated routine looks for new files daily and downloads files via ftp to LASPs LISIRD for access by the NRL2 solar irradiance algorithm. New sunspot files are expected to occur roughly weekly with a latency of approximately 2 weeks. The latency is a result of the organization of the sunspot data files by sunspot group number (i.e. not by calendar date) and the time is takes for a sunspot group to appear and then rotate off the solar disk.

The Solar Irradiance Data Record processing system also operationally access the Mg II index record from LASP’s LISIRD with the same underlying LaTiS time series server. Needs updating. I.e. Which Mg II index, a description, and a reference if not elsewhere

* 1. Programming and Procedural Considerations

Execution speed is rapid and optimization is not needed. Given the facular and sunspot indices, the algorithm calculates the corresponding total and spectral irradiance in less than 2 seconds (on an IMac with Maverick Operating system).

* 1. Quality Assessment and Diagnostics

The quality assurance (QA) process is a subset of both the science analysis and the data quality assurance. The Solar Irradiance Data Record team oversees this process, which involves a careful examination of all solar and proxy data, and judges the reasonableness and quality of the data to be released. The quality assurance takes several different forms based on: 1) the confidence in the calibration and performance of the instrument proving total and spectral solar irradiance observations, 2) comparisons of model output with previous and simultaneous measurements, 3) understanding of the Sun and its variability – an understanding based on solar models and on solar observations at other wavelengths.

The Solar Irradiance Data Record production system supports both automatic and manual diagnostic statistical analyses of the science products. Deviations from expected or predicted values, flagging of anomalous values, and trends of the sunspot blocking function and facular brightening function, as well as final science values, are all incorporated into the assessment of the stability in the final science data products. The Solar Irradiance Data Record team will initially monitor the quality flags in the final science products manually. As the operational implementation of the Solar Irradiance Data Record algorithm matures, the Solar Irradiance Data Record team will move toward automating portions of the quality flag in the final science products; manual monitoring, particularly of the physical representativeness of the facular brightening function and sunspot darkening function, will continue to be necessary to some extent.

A derived relationship between the Mg II index and the F10.7 cm flux (another proxy of chromospheric variability, independent of the Mg II) monitors, identifies, and flags outliers in the Mg II index value that is used to derive the facular brightening function relative to Quiet sun values. In addition, the time series of sunspot number is monitored to screen a value of ‘0’ sunspot area as a “negative” result when the cause is missing station data, or a “positive” result when a value of ‘0’ sunspot area is physically plausible at solar minimum conditions with zero sunspot number.

Statistical analysis of the time series of the mean facular brightening and sunspot darkening functions, and their respective variances, are used to flag potential outliers, alerting the Solar Irradiance Data Record team to investigate the input data sets. The original NRLTSI and NRLSSI models provide robust data sets to define minimum and maximum ranges of the facular brightening and sunspot darkening functions. Statistical monitoring of the time series of the mean modeled solar irradiance (total and spectral), and their respective variances, is also implemented with near real-time comparisons with observations of total and spectral irradiance being made by SORCE, the TSI Transfer Calibration Experiment (TCTE), and the future TSIS instrument suite scheduled to launch to the International Space Station (ISS) in 2017.

Table 6 lists the assumptions in the algorithm. The operational steps to monitor time series of algorithm inputs and the output solar irradiance are also noted.

* 1. Exception Handling

The processing code is not currently robust when it comes to exception handling. Error conditions are unlikely, assuming the processing environment is adequately prepared. Future enhancements will at least ensure that appropriate error messages are returned.

* 1. Algorithm Validation

In addition to the checks on the algorithm performance described in Section 4, the algorithm is validated by comparing the total and solar spectral irradiance times series that it produces with direct observations where available, and with other models of irradiance variability.

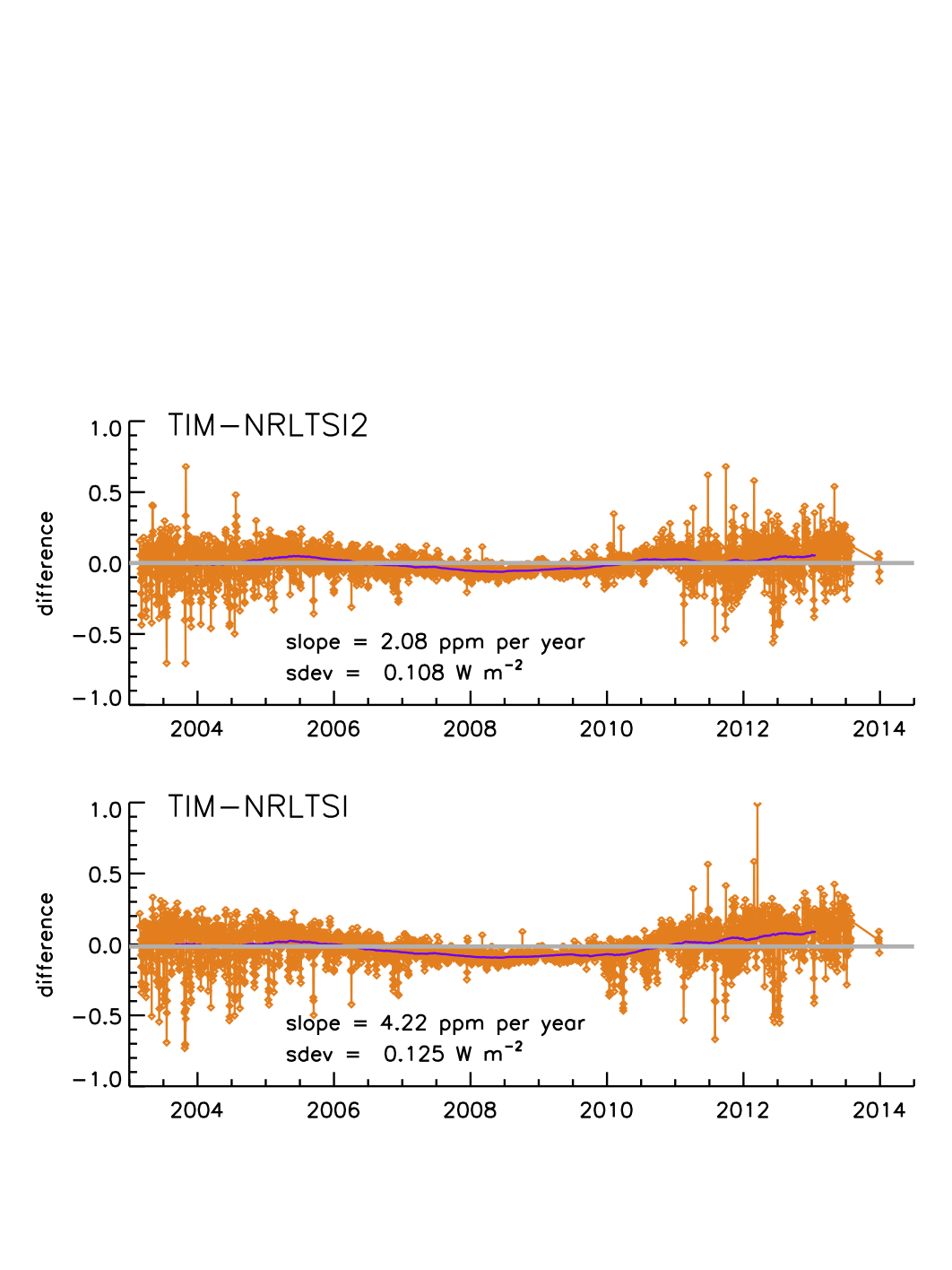
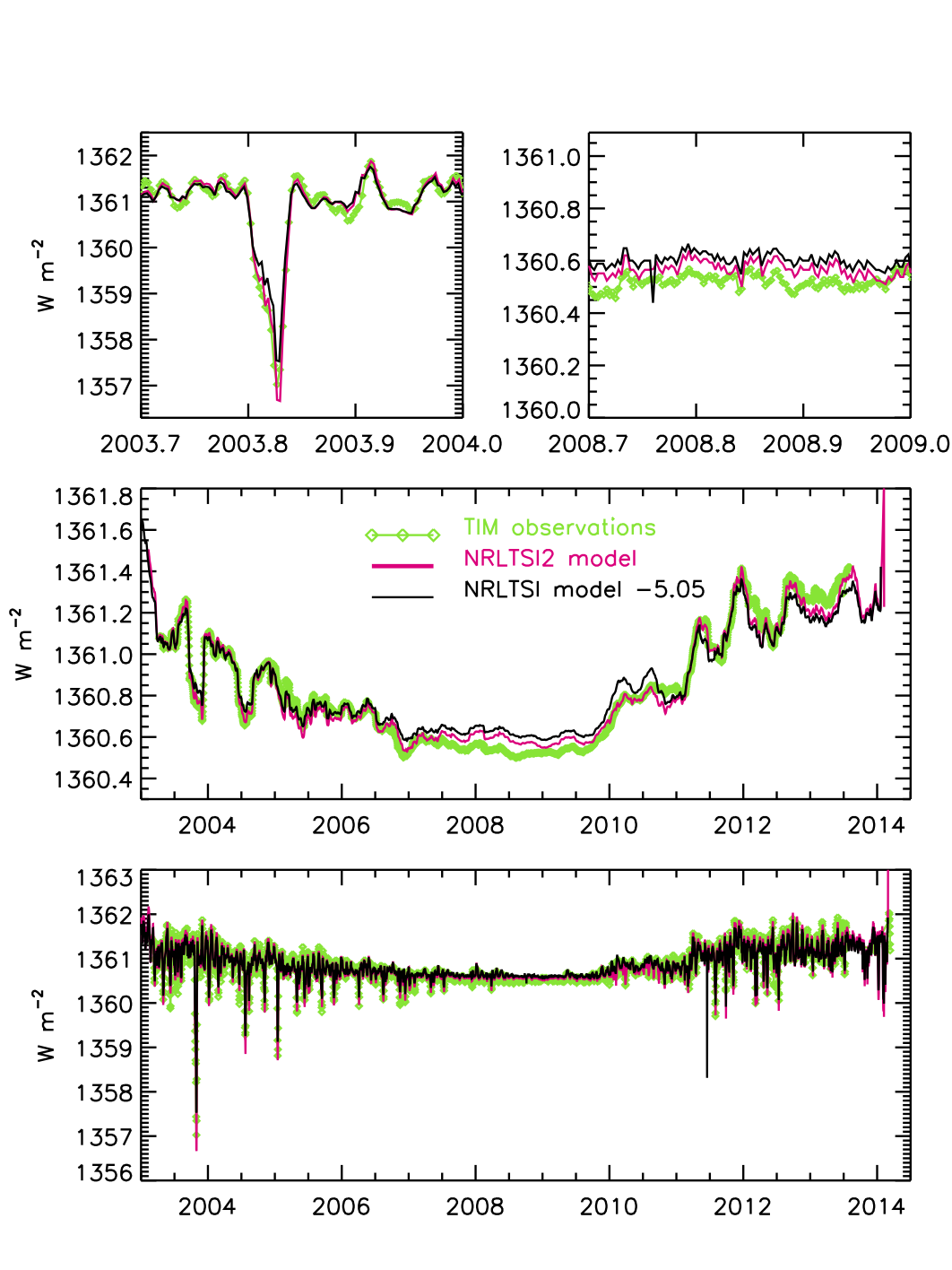
Figures 11 and 12 compare the NRLTSI2 total irradiance time series with the observations made by SORCE/TIM and also with the original NRLTSI model. NRLSTSI2 captures 92% of the observed daily variability (correlation coefficient is 0.96) and provides improved representation of the TIM observations compared with NRLTSI. The NRLTSI2 model-data residuals have a standard deviation of 0.11 W m-2 (compared with 0.13 W m-2 for NRLTSI) and a long-term trend of 2 ppm per year (compared with 4 ppm per year for NRLTSI), which is less than TIM’s 10 ppm long-term stability.

Figure 12. Ratios of NRLTSI2 and NRLTSI models to TIM observations of total solar irradiance variability.

Figure 11. Comparison modeled total solar irradiance time series with observations made by TIM on SORCE, during solar rotation (upper panel) and during the solar cycle (middle and lower panels).

Figure 13 compares detrended time series of solar spectral irradiance in selected broad wavelength bands, calculated by the NRLSSI2 model algorithm, compared with observations made by SORCE SIM and with equivalent time series of the NRLSSI model. The comparisons show spectral irradiance modulation associated with the Sun’s 27-day rotation. On these relatively short (compared with the solar cycle) time scales, instrumental effects in the observations are assumed to be modest compared with true solar irradiance variability.

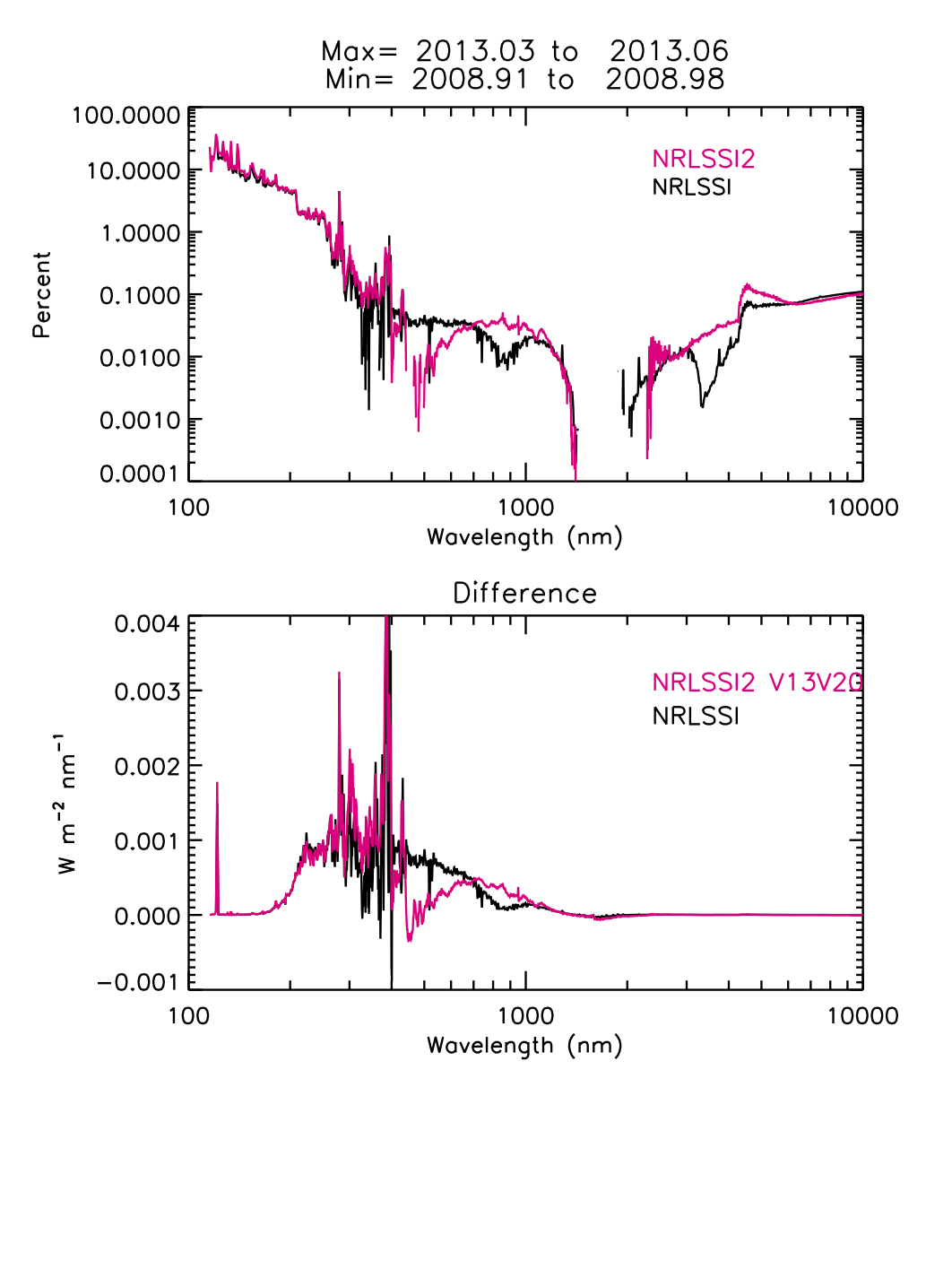
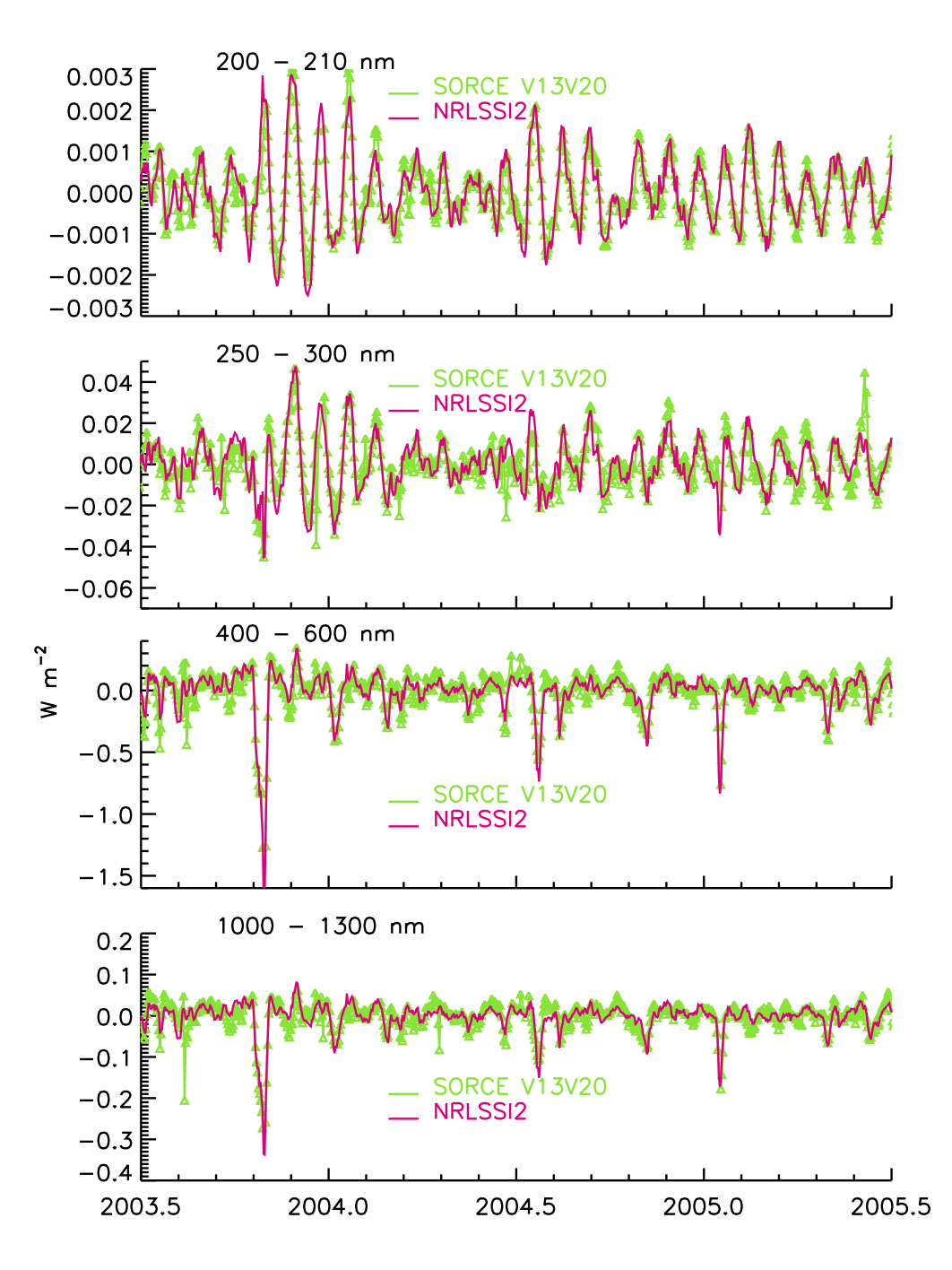
Solar cycle changes in spectral irradiance are presently ambiguous because the only observations are those made by SIM on SORCE whose long-term stability has yet to be validated. Figure 14 shows the solar cycle changes in spectral irradiance estimated by the NRLSSI2 algorithm, compared with the changes estimated by NRLSSI. Compared with NRLSSI, NRLSSI2 is more variable at wavelengths from 300 to 400 nm and less variability at wavelengths from 300 to 600 nm

Figure 14 Solar cycle (24) changes shown as percentages (upper) and in energy units (lower) in the NRLSSI2 model, compared with NRLSSI.

Figure 13 Comparisons of NRLSSI2 solar spectral irradiance in broad wavelength bands with SIM observations. Both the model and observed time series have been detrended; the variations are associated with the Sun’s 27-day rotation.

The Solar Irradiance Data Team will continue to validate the solar irradiance CDR algorithm, with ongoing tests, statistical characterizations and self-consistency checks, and comparison with observations. Planned for the immediate future are separate independent evaluations of the algorithms by both NRL and LASP, and regular comparisons and validation of the algorithm’s operational products.

* 1. Processing Environment and Resources

The processing code is written in IDL and requires version 8.2 or later. It runs on commodity hardware (e.g. 2 cores 8Gb RAM) with modest computation time - no more than one minute to process a day. In principle the code could run on any OS with IDL, though it was developed on Unix based systems. All the processing routines and utility routines are included in the code repository. The only storage concerns are for the final output products.

The primary data access layer is currently provided by LaTiS, an Open Source software framework developed at LASP to provide uniform access to datasets. The sunspot area data from NGDC and the MgII index data from LASP are requested by the IDL processing code from a LaTiS server running at LASP. Other input files are included in the code repository and read directly by the IDL code.

1. Assumptions and Limitations

The assumptions in the algorithm’s theoretical basis are detailed in Section 3. Table 6 summarizes these assumptions and the approach, where possible, to quantifying and monitoring their impacts on the modeled solar irradiance (i.e. validation). The accuracy of the modeled solar irradiance also depends on the inputs (facular brightening and sunspot darkening), but the measurements of the indices used to derive these inputs (the Mg II index and sunspot areas) have uncertainties themselves. In addition, there are uncertainties in the representativeness of the indices to the true (wavelength-dependent) facular brightening and sunspot darkening. Also contributing to the uncertainty in the modeled irradiance, and possibly the easiest to quantify, are statistical uncertainties in the coefficients obtained by multiple regression of the input indices to the measured solar irradiance (total and spectral).

A number of validation studies have been performed, as illustrated in Section 5.5. The uncertainty estimates given in Table 5 are initial first-time values that account for uncertainties in the input indices, and in the multiple regression coefficients and how these translate to equivalent irradiance increments. But these initial estimates do not account for scientific assumptions in the algorithm. Work is ongoing to estimate more realistic uncertainties in the NRLTSI2 and NRLSSI2 modeled solar irradiances, by incorporating better quantification of the algorithm’s assumptions, such as those that Table 6 identifies.

Table 6: Summary of assumptions in the theoretical basis for modeled solar irradiance, model inputs and the potential validation approaches. Particular validation approaches that can be monitored over time (i.e. statistical) to provide an estimate in the uncertainty in the modeled solar irradiance are labeled ‘*Operational’*.

|  |  |
| --- | --- |
| **Assumptions & Scientific Support/Citation(s)** | **Validation Approach** |
| Adopted value of the Quiet Sun (total and spectral) is invariant (Lean et al., 1998). | Comparison to irradiance measurements at solar minimum conditions. |
| Faculae brightening and sunspot darkening are the only modulators of *contemporary* solar irradiance and their respective impacts on irradiance are represented by linear adjustments of baseline, Quiet Sun, conditions (Lean et al., 1998, 2005; Lean, 2000; Lean and Woods, 2010). | The derived facular and sunspot functions are imperfect indicators of the sunspot darkening and facular brightening sources. Small ‘epsilon’ regression coefficients correct for these differences.   * ***Operational***: Monitor the correlation between contemporary modeled and measured solar irradiance, and the standard deviation of residuals between modeled and measured solar irradiance. * Future improvements in the facular and sunspot functions may enable elimination of the small, but non-zero, correction coefficients. |
| The background, facular brightening at *historical*, longer-term time scales (decade, centennial) is speculated (Lean et al., 1992, 1995, 2000, 2002, 2005; Hoyt and Schatten, 1993; Tapping et al., 2007; Radick et al, 1998; Hall and Lockwood, 2004; Wang et al., 2005) | A plausible magnitude of facular brightening is simulated from a flux-transport model that simulates eruption, transport, and accumulation of magnetic flux from Maunder Minimum to present as a function proportional to sunspot number.   * The effect is 27% lower than estimated in original NRLSSI model. * Monitor circumstantial evidence of facular impacts on irradiance trends:   + Reduction in measured TSI corresponding with disappearance of faculae.   + Variations in chromospheric activity of Sun-like stars.   + Inferences with cosmogenic and geomagnetic indices.   + Inferences with changes in solar structure. |
| The sunspot darkening function can be computed using sunspot areas, heliocentric locations, and number of individual sunspot regions (Allen, 1979; Foukal, 1981; Lean et al., 1998; Brandt, Stix and Weinhart, 1994).   * The relative contrast of sunspots to Quiet Sun irradiance is known with these assumptions:   + The Center-to-limb variance is independent of wavelength.   + Sunspot contrast is independent of position on solar disk and has an experimental bolometric (integrated) value of 0.32 with a dependency on sunspot area.   + Sunspot darkening function is the average of all sunspot areas and locations over the day. | The solar rotation axis, ecliptic plane (beta angle) are known throughout the year and are used to adjust the projection of sunspot areas to the direction of the Earth, as it orbits the Sun.   * Sunspot areas prior to 1976 are corrected for a systematic 20% high bias (measurement error). * ***Operational***: Monitor mean and standard deviation of time series of sunspot darkening function. * ***Operational***: Implement quality flags for the sunspot blocking function. Flag:   + Missing station data   + Duplicate records   + Larger (or smaller) than expected variability. * ***Operational***: Monitor relationship between sunspot areas and sunspot number. If sunspot number is zero, a physically plausible result of sunspot area is 0 (i.e. “positive” result). If sunspot number is non-zero, a physically implausible result is a sunspot area of 0 (i.e.“negative” result). |
| The facular brightening function is a linear function of a flux “proxy” of facular brightening (Lean et al., 1998).   * The Mg II index is a proxy for chromospheric variability, which is an extension of photospheric faculae (Snow et al., 2005). * The Mg II index is (relatively) free of instrumental sensitivity (drifts). | Faculae are poorly observed in solar imagery and inadequately specified.   * ***Operational***: Monitor mean and standard deviation of time series of facular brightening function.   + For single, and multiple (i.e. overlapping in time) instruments. * ***Operational***: Monitor relationship between time series of Mg II index and the F10.7 cm flux. The relationship is expected to be consistent; a deviation is expected to be indicative of outlier in the Mg II record. * ***Operational***: Implement quality flags for facular brightening based on Mg II record. Flag:   + Time gaps   + Larger (or smaller) than expected variability.   + Outliers (when compared to F10.7 cm flux). * Future, reliable and quantitative, observations can be used to define statistical definitions of spectral irradiance changes due to faculae. |
| SORCE TIM measurements from 2003 to 2014 are the total solar irradiance standard used to compute scaling coefficients of facular brightening and sunspot darkening for NRLTSI2 using a multiple linear regression technique.   * TSI of Quiet Sun is 1360.8 ± 0.5 Wm-2 (Kopp and Lean, 2011), based on SORCE TIM measured irradiance at solar minimum. | ***Operational***: Monitor the correlation between contemporary modeled and measured total solar irradiance, and the standard deviation of residuals between modeled and measured total solar irradiance. |
| SORCE SIM measurements contain instrumental trends (Lean and Deland, 2012).   * Prior to application of multiple linear regression technique to observations, the SORCE SIM data is detrended (with 81 day running mean). | Compare wavelength-dependent scaling coefficients of facular brightening and sunspot darkening to their respective theoretical contrasts (ratio of emission to quiescent solar atmosphere).  For λ > 295 nm: Use detrended TSI observations to determine the ratio of SSI multiple regression coefficients from direct observations to the detrended observations (i.e. an adjustment factor).  For λ < 295 nm: Estimate adjustment to detrended multiple regression coefficients using Ca K time series; Ca K is a proxy of chromospheric variability independent of the Mg II index. |

The most probable cause of the algorithm generating incorrect irradiance values lies with the determination of the facular brightening and sunspot darkening indices, which rely on ground- and space-based observations of sunspot active regions and global facular brightness. The accuracy and precision of these ground-based observations is essentially unknown, and spurious inputs could produce unrealistic irradiance values. The algorithms flags input values that are deemed implausible, outside the range of current observed values.

Planned for future versions of the algorithms are more sophisticated near-real time validation of the sunspot darkening and facular brightening inputs that will aid in securing a more robust algorithm. e.g., comparisons of simultaneous calculations from independent databases.

* 1. Algorithm Performance

There are no assumptions made concerning algorithm performance. The algorithm is designed to compute modeled solar irradiance, using basic algebra, over a user-defined time range, and this input time range can be of arbitrary length. The algorithm is free of matrix inversions, parameter extrapolations, or interpolations, and the execution is rapid (insert running time as a function to the number of steps, or storage locations) and repeatable.

* 1. Sensor Performance

The solar irradiance reconstructions that this C-ATBD describes compliment the direct measurements made of total and spectral solar irradiance by the Total Solar and Spectral Solar Irradiance Sensor (TSIS) instrument, documented in the TSIS ATBD (**ref)**. The TSIS ATBD describes the algorithms used to produce all data levels of solar and spectral irradiance for the TSIS instrument complement, which consists of the Total Irradiance Monitor (TIM) and Spectral Irradiance Monitor (SIM). The TSIS ATBD also describes the predicted science and housekeeping operation modes, measurement error budgets, and the plan to monitor and correct for instrument degradation.

For solar irradiance, variations of less than 0.1% per decade are typical of the kinds of signals that must be extracted from “noisy” time-series measurements. The calibration approach adopted for TSIS characterizes the flight instrument as an “absolute sensor”. This involves characterizing each term in the measurement equation and tabulating a list of individual uncertainties and root sum square errors for overall measurement uncertainty.

***Accuracy and Long-term Stability of the TSIS Instrument Complement***

The TSIS TIM instrument is estimated to be three times more accurate than the SORCE TIM due to engineering advances in the optical and electrical sensors and to the end-to-end validation of the radiometer at the TSI Radiometer Facility (TRF) at the Laboratory for Atmospheric and Space Physics (LASP).

Lessons learned from the first-ever measurements of spectral solar irradiance made by the SORCE SIM have been incorporated by TSIS SIM to meet the measurement requirements. The specific TSIS SIM capabilities over SORCE SIM include reduced uncertainties in the prism degradation correction to meet long-term stability requirements, improved noise characteristics of the electrical substitution radiometer (ESR) and photodiode detectors to meet the measurement precision requirement, and improved absolute accuracy through pre-launch calibration using the novel, Spectral Radiometer Facility (SRF) at LASP.

The TSIS TIM will measure 4x daily total solar irradiance with an absolute accuracy of 100 ppm and a relative accuracy of 10 ppm. The TSIS SIM will measure 2x daily solar spectral irradiance at variable resolution from 200-2400 nm with an absolute accuracy of 0.2% (2000 ppm), a relative accuracy of 0.02% (200 ppm), and with long-term relative stability of 0.05% per year (for wavelengths shortward of 400 nm) and 0.01% per year for wavelengths longward of 400 nm.

SORCE SIM data is currently being reanalyzed and its uncertainty estimates may change. It has been suggested that the SORCE SIM time series has not been fully corrected for instrument degradation and that the long-term trends are not solely of solar origin (*Lean and Deland*, 2011). To reduce the dependence of the modeled solar spectral irradiance to these potential instrumental trends, the respective coefficients for sunspot darkening and facular brightening are derived from a multiple linear regression technique applied to the sunspot area and Mg II index data, respectively, and the *detrended* (i.e. subtraction of an 81-day running mean) SORCE SIM observations (as described in Section 3.4).

***Degradation Monitoring and Correction of the TSIS Instrument Complement***

The sensitivities of all instruments on TSIS are assumed to degrade as the mission progresses and solar exposure accumulates. There are general assumptions that the degradation will be small, that the sensitivity will monotonically decrease with time, and that a primary cause of the decreased sensitivity is related to the exposure to the harsh radiation environment from the Sun. However, there is no guarantee these assumptions will be met and other changes that are strictly time-dependent, or aging effects, must also be considered. The possibility that instrument sensitivity may increase cannot be ruled out, and the degradation analysis does not preclude this condition. That is, the analysis is open to the possibility that aging effects improve sensitivity, or that radiation and exposure may in fact “scrub” surfaces and improve their throughput.

Exposure and time-dependent degradation is a challenging problem and will require refinements throughout the mission as well as considerable analyses effort by the TSIS instrument scientists. In addition, any correction parameterization used initially may need evaluation and modification during the mission. Unexpected changes to the thermal stability of the spacecraft environment may require offsets in the analysis of certain data, and electronics and detector functionality can be impacted by energetic particles in major solar storms. Such impacts may require discontinuous changes in the science product rather than parameterized functions.

The TSIS instrument science team will examine pre-launch characterizations and in-flight calibrations to derive appropriate degradation corrections. The technique used to understand instrument degradation for the TSIS TIM and SIM instruments is to have completely independent instrument channels, to use each channel with a varying duty cycle, and then to compare their observations of the Sun. The degradation in the instruments is assumed to be primarily dependent on the exposure of the optics and detectors to solar radiation. If the exposure times are dramatically different, say one to one hundred, the ratio of the measurements will change with time. With the assumption that exposure-dependent degradation will proceed proportionally faster for the normal channel, an exposure-dependent model of degradation is developed. The TSIS ATBD outlines one proposed model for monitoring and correcting each TSIS SIM observation for degradation.

1. Future Enhancements

Planned in the immediate future are more sophisticated determinations, and near-real time validation, of the sunspot darkening and facular brightening inputs. Improved algorithm inputs will aid in securing more robust algorithm output. Also planned are improved uncertainty estimates, taking into account the assumptions in the algorithm formulation. On longer time scales, the algorithm may be revised and updated with a new formulation framework and coefficients, to reflect improved understanding of solar irradiance absolute scale and variability from ongoing analysis and recalibration of existing irradiance (and other) datasets, as well as the availably of reprocessed and new irradiance databases, such as from TCTE and TSIS observations.

Because solar irradiance is an essential, universal input for models of the terrestrial environment, we also anticipate provision of new irradiance products to accommodate the evolving needs of the user community, which we plan to monitor closely.

* 1. Enhancement 1: Improved Sunspot Darkening Index

Preliminary comparisons of the NRLTSI2 and NRLSSI2 models with observations suggest that the model overestimates (slightly) the irradiance reduction when the sunspot darkening is large and underestimates it (slightly) when the sunspot darkening is small. This is visually evident in the comparison of the NRLTSI2 model with the TIM observations during the largest rotational modulation that SORCE has observed (in October 2003), shown in Figure 11 (upper left plot). It is also evident in the comparison of the NRLSSI2 irradiance in broad wavelength bands with SIM observations, shown in Figure 13.

A likely reason for these differences is the less than optimal parameterization in the algorithm of the sunspot contrast with area, which is given in Section as

Future work is planned to investigate this dependence of sunspot contrast on area by calculating different sunspot darkening functions and evaluating their performance in the model formulation.

In the current algorithm, the sunspot darkening is calculated as the (unweighted) average of sunspot region information from ten different ground stations that contribute to the Air Force SOON program. We will investigate the robustness (repeatability and relaibility) of the individual stations and consider a weighted average of their observations (also taking into account the local time of their observations, which alters the reported sunspot location on the disk) depending on the outcome of statistical comparisons. As well, different stations have different longevity within the record, with some stations terminating and some commencing at different times. We will investigate how the use of different combinations of observing stations affect the computed sunspot darkening function and how these differences translate into uncertainties in the input solar darkening index and hence in the solar irradiance that the algorithm outputs.

New, improved uncertainties in the sunspot darkening input to the algorithm will be estimated based on the results of the above analyses, as well as from comparisons with independent sunspot darkening calculations, such as in STARA (<http://www.nso.edu/staff/fwatson/STARA>) and from the Debrecen (<http://fenyi.solarobs.unideb.hu/DPD/>).

* 1. Enhancement 2: Improved Facular Brightening Index

Because the solar irradiance climate data record algorithm is formulated from multiple regression of the facular brightening and sunspot darkening indices with the TIM observations, changes in either index necessarily affect the proportion of irradiance variability that the model ascribes to each. An improved sunspot darkening index will therefore enable better assessment of limitations in the current facular index. We will investigate possible improvements to the facular index by analyzing the residuals of the observed irradiance and improved sunspot darkening indices (obtained from Enhancement 1, described above).

More problematic and challenging is securing a facular index that has robust long-term (solar cycle time scale) stability. There are distinct and reported differences among various Mg indices produced by different groups. In particular, changes from the successive 1996 to 2008 solar minima are not consistent (within current understanding) among various solar proxy indices (Snow et al., 2014). Analysis and study by the solar irradiance community is ongoing to secure the most reliable facular index possible. We will test the effect of new indices in the model formulation and algorithm, as well as compare different indices to better quantify irradiance uncertainties arising from this index.

As with the sunspot darkening analysis in Enhancement 1, future analysis of various facular brightening indices will enable improved uncertainties for this input to the solar irradiance algorithm.

* 1. Enhancement 3: Improved Uncertainty Estimates

Table 5 provides initial estimates of time-independent uncertainties in solar irradiance that the algorithm calculates, at selected wavelengths. In reality, the uncertainties are expected to vary with solar activity, which dramatically alters the magnitude of the sunspot darkening and facular brightening indices, as well as with wavelength, since the sunspot and facular contrasts have strong wavelength dependences. A future enhancement will expand and improve the uncertainty estimates in Table 5., taking into account the improved quantitate estimates of uncertainties in solar irradiance arising from the sunspot darkening and facular brightening indices (described in Enhancements 1 and 2).

Most importantly, we plan to Investigate and attempt to quantify uncertainties arising from assumptions in the model formulation, such as those that table 6 summarizes.

* 1. Enhancement 4: Improved Model Formulation

The solar irradiance algorithm uses coefficients that require three databases for their determination: a time series of observed solar irradiance, and time series of the facular brightening and sunspot darkening indices. Significant (or even modest) changes in any of these time series may result in altered algorithm coefficients and warrant a new algorithm version. The sunspot and facular time series may change as a result of Enhancements 1 and 2, and the solar irradiance dataset may change as a result of reprocessing existing observations or by the addition of new observations.

It is possible that the current model formulation framework may also change in the longer-term future, to better reflect new understanding of the causes of solar irradiance variability, and taking into account the assumptions in Table 6. For example, the introduction of additional terms to the regression analysis would result in a new model formulation and CDR algorithm.

* 1. Enhancement 5: Additional Irradiance Products

The current solar irradiance climate data record algorithm produces the solar spectral irradiance in 1 nm (or broader) spectral bands on a 1 nm (or broader) wavelength grid. However, fundamental, line-by-line calculations of atmospheric heating rates require the input solar spectral irradiance at higher spectral resolution and on a wavenumber grid. So a possible future additional irradiance product is might be solar irradiance on a higher resolution wavenumber grid.

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TSIS ATBD

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1. Acronyms and Abbreviations

|  |  |
| --- | --- |
| Acronym or Abbreviation | Meaning |
| λ | Lambda; wavelength (nm) |
| μ | Heliographic position |
| AA index | Magnetic activity index |
| ACRIM | Active Cavity Radiometer Irradiance Monitor |
| C-ATBD | Climate Algorithm Theoretical Basis Document |
| Ca K | Calcium Potassium (emission line of the Sun) |
| CDR | Climate Data Record |
| F(t) | Facular Brightening Function (time-dependent) |
| F10.7 | Solar flux at 10.7 cm |
| GOME | Global Ozone Monitoring Experiment |
| GOME-2 | Global Ozone Monitoring Experiment 2 |
| I(t) | Solar Spectral Irradiance (time-dependent) |
| IQ | Solar Spectral Irradiance of the Quiet Sun |
| ISS | International Space Station |
| LASP | Laboratory for Atmospheric and Space Physics |
| LISIRD | LASP Interactive Solar Irradiance Data Center |
| m | meter |
| Mg II | Magnesium II index |
| MDI | Michelson Doppler Imager |
| NCDC | National Climatic Data Center |
| netCDF4 | Network Common Data Format |
| NOAA | National Oceanographic and Atmospheric Administration |
| NGDC | NOAA National Geophysical Data Center |
| NRL | Naval Research Laboratory |
| NRLSSI | Naval Research Laboratory Solar Spectral Irradiance model (original) |
| NRLSSI2 | Naval Research Laboratory Solar Spectral Irradiance model (version 2) |
| NRLTSI | Naval Research Laboratory Total Solar Irradiance model (original) |
| NRLTSI2 | Naval Research Laboratory Total Solar Irradiance model (version 2) |
| nm | nanometer |
| NOAA | National Oceanic and Atmospheric Administration |
| PMOD | Physikalisch-Meteorologisches Observatorium Davos |
| ppm | Part per million |
| QA | Quality Assurance (Analysis) |
| R | Radiance |
| S(t) | Sunspot Darkening Function (time-dependent) |
| SCIAMACHY | SCanning Imaging Absorption SpectroMeter for Atmospheric CHartographY |
| SIM | Spectral Irradiance Monitor |
| SMM | Solar Maximum Mission |
| SOHO | Solar and Heliospheric Observatory |
| SOLSTICE | Solar Stellar Intercomparison Experiment |
| SOON | Solar Observing Optical Network (US Air Force) |
| SORCE | Solar Radiation and Climate Experiment |
| SRF | Spectral Radiometer Facility (LASP) |
| SSD | Space Science Division (Naval Research Laboratory, Washington, DC) |
| SSI | Solar Spectral Irradiance |
| T(t) | Total Solar Irradiance (time-dependent) |
| TQ | Total Solar Irradiance of the Quiet Sun |
| TCTE | TSI Transfer Calibration Experiment |
| TIM | Total Irradiance Monitor |
| TRF | TSI Radiometer Facility (LASP) |
| TSI | Total Solar Irradiance |
| TSIS | Total and Spectral Solar Irradiance Sensor |
| UARS | Upper Atmosphere Research Satellite |
| UV | ultraviolet |
| W | Watt |