








Thermodynamic Structure of the Solar Corona: Tomographic Reconstructions and MHD Modeling

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Abstract Observational tools play an essential role in advancing our understanding of the physics of the solar corona. They provide validation data to improve three-dimensional (3D) magnetohydrodynamic (MHD) models of the

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solar atmosphere, key to enhance their space weather forecasting capabilities. Solar rotational tomography (SRT) is currently the sole observational technique that provides an empirical 3D description of fundamental plasma parameters of the solar corona at a global scale. Based on EUV data of space borne instruments, SRT allows constructing 3D maps of the coronal electron density and temperature at heliocentric heights below 1.25 R_{sun}. We carry out a study of the corona combining tomographic reconstructions with MHD simulations of the latest version of the Alfvén Wave Solar Model (AWSoM) of the Space Weather Modeling Framework (SWMF). Target rotations were selected from the solar minimum between solar cycles (SC) 23 and 24 and the current declining phase of solar cycle 24. Tomographic reconstructions and results of the model are analyzed in distinct coronal magnetic structures. We study magnetically closed structures associated with the equatorial streamer belt, as well as magnetically open regions enclosing it. We report on the tomographic results in the different structures, their implications for the physics of the solar corona, and the capability of the AWSoM model to reproduce the tomographic reconstructions in different regions.

Keywords: Solar Cycle, Observations; Corona,E; Corona, Structures

1. Introduction

2. Methodology

2.1. DENT reconstructions

2.2. AWSoM simulations

2.3. Tracing results along fieldlines

3. Results

3.1. Tomographic Results

A) Carrington (Lat/Lon) maps of the DENT products Ne and Te at selected heights 1.025, 1.105, 1.205 R_{sun}, over plotting the Open/Closed (O/C) boundary from the AWSoM model. This plot gives initial CONTEXT.

B) Carrington maps of the Lat/Lon location at 1.105 Rs of the field lines I (in blue), II (in red), and III (in green). This shows in a very graphical way which regions of the corona are probed by the three classes of selected field lines.

C) A plot of the average DENT Ne(r) along all field lines of type I (in solid-blue), of type II (in solid-red and of type III (in solid-green). Over plotted in the same color code but in dashed-line style the AWSoM results along the same three statistical sample of field-lines. In particular, this shows the global 3D agreement between AWSoM and DENT in distinct magnetic structures and carefully tracing along field lines

D) Same as C) but for Te.

C) and D) will be complemented with companion histograms showing in a statistical fashion how AWSoM and DGMT results compare.

E) As the two rotations are strongly axis-symmetric we want to show the longitude-averaged results of DGMT at a low height, say 1.105 R_{sun}, along with AWSoM results. In the attached plot we show the DGMT and AWSoM longitude-averaged results at 1.105 R_{sun} in the top two panels (one for each rotation). The vertical lines indicate the longitude-averaged latitudes of the North and South OC boundary. The LOW panels show the longitude-averaged AWSoM wind speed V_r at 6 R_{sun}, where all field lines are open. The HCS location is indicated by the minimum of the speed curve. Everything to the South of the HCS maps down to the Southern OPEN region in the top panels, and the same for the Northern hemisphere. It is pretty obvious the ANTI-CORRELATION between the tomographic Ne at low heights (coronal source region of the wind) and the V_r.

This plot (E) has never been shown before and I believe it is a very nice result. There are of course previous studies on this, but this would be the first time this is evaluated with tomography. It is not the terminal speed, but it already shows the fast/slow components. We can show this, or we can go for more. Diego can QUANTIFY the correlation by carefully mapping all open field lines at 6 R_{sun} down to the DGMT region. He needs to code a bit more to achieve this goal, but we believe it is worth. With that tool done Diego can explore other terminal values of quantities of the AWSoM model and the DGMT results below (expansion factor maybe?). This will be nice to do also for the July-2-2019 Eclipse rotation that we will analyze after this paper.

3.2. Tomography and Models

Same as A but for AWSoM model. Same as B but for AWSoM model

4. Discussion