

Thermodynamic Structure of the Solar Corona: Tomographic Reconstructions and MHD Modeling

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Abstract Observational techniques play an essential role in advancing our understanding of the physics of the solar corona. They provide validation data

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for three-dimensional (3D) magnetohydrodynamic (MHD) models of the solar atmosphere, key to improve their space weather forecasting capabilities. Solar rotational tomography (SRT) is currently the sole observational technique that provides an empirical 3D description of some of the fundamental plasma parameters of the solar corona at a global scale. Based on EUV data of space borne instruments, SRT allows constructing 3D maps of the coronal electron density and temperature at heliocentric heights below 1.25 R_{sun}. We carry out a study of the corona combining tomographic reconstructions with MHD simulations using the latest version of the Alfvén Wave Solar Model (AWSoM) of the Space Weather Modeling Framework (SWMF). Target rotations were selected from the solar minimum between solar cycles (SC) 23 and 24 and the current declining phase of solar cycle 24. Tomographic reconstructions and results of the model are analysed in distinct coronal magnetic structures. We study magnetically closed structures associated with the equatorial streamer belt, as well as magnetically open regions enclosing it. We report on the tomographic results in the different structures, their implications for the physics of the solar corona, and the current capability of the AWSoM model to reproduce the tomographic reconstructions in different regions.

Keywords: Solar Cycle, Observations; Corona,E; Corona, Structures

1. Introduction

Being the place where the solar atmosphere is heated and the solar wind accelerated, and where impulsive events such as solar flares and coronal mass ejections are energized, observation and modeling of the solar corona are tasks of great relevance to improve our understanding of the Sun-Earth environment. To advance our knowledge of the physics of the solar corona, as well as to improve its three-dimensional (3D) models, information derived from observational data plays a key role. Solar rotational tomography is currently the sole observational technique able to provide a quantitative empirical description of the 3D distribution of some fundamental plasma parameters of the solar corona at a global scale.

NOTE: Here we will put in a couple of paragraphs summarizing what SRT and DENT are, including references and defining the acronyms in the process. We should refer to minima in particular. We should mention the instruments and missions we use as data for DENT, and define their acronyms too. Also, a couple of paragraphs on the AWSoM model and the SWMF, including references and defining the acronyms in the process.

2. Methodology

2.1. DENT reconstructions

NOTE: This section will summarize DENT's main aspects and equations, only those needed to understand the paper. Nothing not needed to read this paper is

provided. Also, no derivation of equations are provided. References are given so that the reader can dig deeper if so desires. Points to cover are:

- Selected targets, instruments used. One paragraph.
- Characteristics of the data (specific time series, dates, bands, etc.), SolarSoft software version used for their processing. One paragraph.
- Technical aspects of DENT (grid, regularization, off-limb data is only used) and how these differ from previous versions (references). One paragraph.
- What is FBE, what is LDEM, what is R. Why we care about these three quantities and what is their physical meaning should be clear. Then the final products N_m and T_m , and what is their physical meaning.

To perform the comparison two specific rotations were selected, CR-2082 (2009, 05 April through 03 May), a deep minimum period between solar cycles (SCs) 23 and 24 characterized by virtually no ARs, and CR-2208 (2018, 02 September through 29 September), a rotation during the early declining phase of SC 24.

To determine the thermodynamical structure of both rotations, the DENT technique was applied to respective time series of EUV images covering each first half rotation. To study CR-2082 and CR-2208 data taken by the EUVI-B/STEREO and AIA/SDO instruments was used, respectively.

EUV images are affected by instrumental stray light contamination, which can be corrected by deconvolution of the point spread function (PSF). While in the case of AIA images there is no reliable determination of its PSF, it has been well determined for EUVI-B by Shearer *et al.* (2012). For this article, all EUVI images were stray-light corrected. For the details of impact in N_m and T_m , refer to Lloveras *et al.* (2017).

In DENT, the low corona in the height range 1.0 to 1.25 R_\odot is discretized in a spherical computational grid. The size of the tomographic cell (or voxel) is set to 0.01 R_\odot in the radial direction, while is set 2° in both the latitudinal and longitudinal directions. We carried out a reconstruction using half of a solar rotation blocking the lines-of-sight coming from the disk (i.e. only using off-limb data).

Using the high cadence of data available, one image every 10 minutes were used. The selected images were averaged in 6-hour intervals which, given the angular resolution, is the cadence needed to fully constrain the inversion problem, for a total of about 55 images to cover a half solar rotation. For both rotations, the first 14 days of data were used.

DONDE QUEDA MEJOR COMENTAR ESTA MEJORA? Unlike in previous works, a 3D regularization scheme was applied, thus spurious artifacts (induced by the previous longitude latitude regularization scheme) were minimized in the first two heights of the tomographic grid.

We refer the reader to Frazin, Vásquez, and Kamalabadi (2009) for a detailed description of the technique, and to Vásquez (2016) for a recent review on all published work based on it. Here we summarize the key points. The technique involves two consecutive procedures.

In a first step, the time series of EUV images is used to solve a SRT problem, for each EUV band independently. As a result, the 3D distribution of the so called

filter band emissivity (FBE) is determined for each band separately. The FBE, an emissivity-type quantity, is defined as the wavelength integral of the coronal EUV spectral emissivity and the telescope's passband function of each EUV channel. Line-of-sight (LOS) integration of the FBE provides synthetic images that can be quantitatively compared to the real data in the time series. To find the FBE, the tomographic problem is posed a global optimization problem in which the quadratic norm of the difference between all pairs of synthetic and real images is minimized. Spatial regularization terms are also included in the objective function as to minimize high frequency spurious artifacts that are due to the specific characteristics of the SRT sparse projection matrix (Frazin, Vásquez, and Kamalabadi, 2009; Frazin, 2000). Due to optical depth issues that may affect LOSs passing very close to the limb, and to the decay of intensity with radii, the 3D maps produced by DENT analysis typically cover the height range 1.02 to 1.23 R_{\odot} (Frazin, Vásquez, and Kamalabadi, 2009).

COMO INTRODUCO EL NIVEL DE REGULARIZACION? For each instrument, the DENT analysis involves a cross-validation study to determine the optimal level of regularization.

In a second step, the FBE values obtained for all bands in each voxel are used to constrain the determination of a local differential emission measure (LDEM) distribution, which describes the temperature distribution of the electron plasma contained in each individual tomographic grid voxel. Specifically, at each tomographic voxel i , the FBE of the band k is related to the LDEM of the voxel according to

$$\text{FBE}_i^{(k)} = \int dT \text{LDEM}_i(T) \text{TRF}^{(k)}(T), \quad k = 1, \dots, K \quad (1)$$

Due to unresolved coronal dynamics, tomographic reconstructions exhibit negative values of the reconstructed FBE, or zero when the solution is constrained to positive values (Frazin, 2000; Frazin, Vásquez, and Kamalabadi, 2009). These non-reconstructed voxels are indicated as black voxels in the Carington maps of the reconstructed DENT results in Section 3.

Once the LDEM is determined at each voxel, the average squared electron density $\langle N_e^2 \rangle$ and the electron mean temperature T_m in the voxel can be computed by taking its zeroth and first moments over temperature. More specifically, at the i -th voxel,

$$N_{m,i}^2 = \langle N_e^2 \rangle_i = \int dT \text{LDEM}_i(T), \quad (2)$$

$$T_{m,i} = \langle T_e \rangle_i = \frac{1}{\langle N_e^2 \rangle_i} \int dT T \text{LDEM}_i(T), \quad (3)$$

We define next a measure of the accuracy of the LDEM model to predict the tomographic FBEs in each voxel, as

$$R_i \equiv (1/K) \sum_{k=1}^K \left| 1 - \text{FBE}_{i,\text{syn}}^{(k)} / \text{FBE}_{i,\text{tom}}^{(k)} \right|, \quad (4)$$

being the average relative difference between the tomographic and the synthetic FBEs. The final product of DEMENT is in the form of 3D maps of the electron density and mean temperature. These relationships were derived in detail in Frazin, Vásquez, and Kamalabadi 2009 (see Appendix C).

2.2. AWSoM simulations

NOTE: This section will summarize the AWSoM model. After a first version is written we will surely require some input from the Michigan friends. Points to cover:

- Data input. Link with the data input for DEMENT in terms of dates, etc. One paragraph.
- Main physical and technical aspects (grid in particular) of the model. One paragraph.
- Explain the artificially extended TR. One paragraph.
- What are the products of the model we care about in this paper. One paragraph.

In the transition region (TR), the plasma heats up and becomes less dense by several orders of magnitude over a short distance. Modeling this region realistically has a very high computational cost. For this reason, the model uses an artificially extended TR about $\approx 0.05 R_\odot$ thick, as well as artificially large values for the electronic density and temperature as boundary conditions in $r = 1 R_\odot$. Consequently, model results below $\approx 1.05 R_\odot$ were not taken into account in this article.

2.3. Tracing of results along fieldlines

NOTE: This section will explain how the AWSoM magnetic model is used to trace the thermodynamic products of both DEMENT and AWSoM along the fieldlines of AWSoM. This requires covering the following points:

- A first paragraph qualitatively describing the process. This includes mentioning that the AWSoM products are first interpolated into the DEMENT grid.
- A short paragraph describing the numerical method used to trace field lines. We have a nice version of this in your previous paper. At this point the reader should understand that for each field line we have its geometry $\mathbf{r}(s)$ as a function of the distance s along the loop, with $s(1 R_\odot) \equiv 0$.
- Explain very simply how then that for each voxel the middle point of the field line is selected and assigned the DEMENT and AWSoM products of that

cell. At this point the reader should understand that for each field line we have: $Q[r(s)]$, where Q is any given quantity from either the DENT model (specifically $\sqrt{N_m^2}$ and T_m) or the AWSOM model (specifically N_e and T_e).

- Explain fits to the DENT results $\sqrt{N_m^2}(r)$, $T_m(r)$.
- Explain then the classification of field lines.

Aiming determination of the electron density and temperature along individual magnetic field lines, first both the thermodynamic results and the magnetic field obtained with the AWSOM model were interpolated into the DENT grid. Then, the geometry of the field lines is determined by numerical integration of the first order differential equations $dr/B_r = r d\theta/B_\theta = r \sin(\theta) d\phi/B_\phi$, both inwards and outwards, from the specified 3D coordinates of a starting point. In order to evenly sample the whole volume spanned by the DENT reconstructions, one starting point is selected at the center of each tomographic cell at 6 uniformly spaced heights, ranging from 1.025 to 1.225 R_\odot , and every 2° in both latitude and longitude, for a total of 96,000 starting points. Each traced field line is classified as “open” if it intersects the source surface, set in this work at $r = 6 R_\odot$ (where lines become radial), or as “closed” otherwise. In the case of closed magnetic field lines, loops are further classified as “small” and “large” upon their apex height being within or beyond the range of heights studied by DENT, respectively.

At this stage, all DENT and AWSOM products, electron density and mean temperature in particular, can be traced along open and closed magnetic field lines. Once the field line geometry is computed in high spatial resolution, only one sample point per tomographic cell is kept (the median one). To each sample point, the results corresponding to the tomographic voxel where it is located are assigned to it. As a result, for each field line one data point per tomographic cell is obtained. This allows to analyze how density and temperature, coming from both models, vary with height along each individual field line. This approach was firstly used by Huang *et al.* (2012) to study inverted temperature structures in the solar minimum corona, and later on applied by Nuevo *et al.* (2013) to expand that analysis to rotations with different level of activity.

aca estoy haciendo la distincion entre la densidad y temp obtenida por demt y awsom For each open field line, and for each leg of the closed ones, an exponential fit was then applied to the electron density and a linear fit applied to the electron temperature. For DENT results, the data points used were $\sqrt{N_m^2}(r)$ and $T_m(r)$, in case of AWSOM the data points were $N_e(r)$ and $T_e(r)$. The exponential and linear fit equations are described by

$$\sqrt{N_m^2} = N_0 \exp[-(h/\lambda_N) / (r/R_\odot)] \quad (5)$$

$$T_m = T_0 + m h \quad (6)$$

where $h \equiv r - 1 R_\odot$ is the coronal height measured from the photosphere. In the electron density fit, $\lambda_N [R_\odot]$ is the density scale height and $N_0 [\text{cm}^{-3}]$ is the electron density at the footpoint ($h = 0$) of the loop. In the electron temperature fit, $m [\text{MK}/R_\odot]$ is the slope and $T_0 [\text{MK}]$ is the electron temperature at the footpoint of the loop. The slope m is the radial gradient of the electron temperature

along the loop, which we denote as $m = \nabla_r T_m$ hereafter, being $\nabla_r \equiv \mathbf{e}_r \cdot \nabla$ the radial derivative operator, where \mathbf{e}_r is the heliocentric radial unit vector.

In the case of the electron density, the fitted function corresponds to the isothermal hydrostatic equilibrium solution, allowing for variation of the solar gravitational acceleration with height. This choice of function provides a straightforward means to directly evaluate how compatible is the observed coronal thermodynamical state with the hydrostatic solution.

In the present work we focus on lines with increasing and decreasing temperatures in linear form. In other words, we selected those legs where the Pearson correlation coefficient ρ between temperature T and radius r meets $|\rho(T, r)| > 0.5$. (with p-value less than 0.05). In this way, those legs in which the DMT results cannot ensure a noticeable change in temperature between the base and the apex, are discarded.

The linear fit allows characterization of its variation with height by means of a characteristic temperature gradient $\nabla_r T_m$ [MK/R_⊙] along each leg. **En realidad esta regresion lineal utiliza de por si los errores, pero luego tmb los uso para el test de hipotesis, quizas vale la pena ponerlo solo ahi.**

To measure how good the quality of the exponential fit to the electron density and the linear fit to the temperature data are, a chi-square merit function was used (see Section 15.2 Press *et al.* (2002)). The errors were estimated in Lloveras *et al.* (2017), 0.07 MK for T_m and $4.0 \cdot 10^6 \text{ cm}^{-3}$ for N_e and legs with p-value > 0.005 were rejected.

To be selected a leg must meet all following conditions.

- i) The leg must go through at least five tomographic grid cells with reconstructed data, and there must be at least one data point in each third of the range of heights spanned by the leg.
- ii) The value of the Pearson correlation coefficient meets $|\rho(T, r)| > 0.5$.
- iii) The p-value of both the exponential and linear fit must be less than 0.05.

Coronal magnetic structures for which their temperature increases/decreases with height (in the inner corona) have been dubbed as “up”/“down” loops by Huang *et al.* (2012) and Nuevo *et al.* (2013), who first reported their presence by means of DMT studies. As speculated by the authors of those works, loops of type down can be expected if the heating deposition is strongly confined near the coronal base of a magnetic loop. Down loops were first predicted by Serio *et al.* (1981), and later on by Aschwanden and Schrijver (2002). In a recent study, Schiff and Cranmer (2016) reproduced both down and up loops by means of numerical simulations, using a 1D steady-state model and considering time-averaged heating rates.

To characterize the global thermodynamic state of the inner solar corona in distinct magnetic structures, the DMT and AWSOM results were traced along the magnetic field lines of the latter model. Based on the geometry and size of the loops, as well as on their thermodynamical properties, their legs were classified in four different types in this work:

Hice las definiciones con gradiente de temperatura para que quede mas claro, pero en realidad uso el pearson. NOTE: Charlemos los tres en persona de esto.

- Type 0 - Small (meaning with apex within the DGMT range of heights) closed field lines in the latitudes range $[-30^\circ, +30^\circ]$ and $\nabla_r T_m < 0$.
- Type I - Small closed field lines in all latitudes, with $\nabla_r T_m > 0$.
- Type II - Large closed field lines (meaning their apex surpasses the DGMT range of heights) in the latitudes range $[-90^\circ, -30^\circ] \cup [+30^\circ, +90^\circ]$ and $\nabla_r T_m > 0$.
- Type III - Open field lines (meaning with apex out the DGMT range of heights) in the latitudes range $[-90^\circ, -60^\circ] \cup [+60^\circ, +90^\circ]$ and $\nabla_r T_m > 0$.

In Section 3, the results of both the DGMT and AWSOM models in the distinct magnetic structures are statistically analysed.

2.4. Energy input flux

NOTE: This section will summarize how loop integrated energy quantities are computed from the DGMT products and the AWSOM magnetic model alone. The summary will not include any derivation and will refer to Cecilia's paper. Only the results that are relevant are to be explicated. Start by explaining the model and assumptions, and end up summarizing the loop-integrated products (radiative and conductive fluxes). At the end the reader should understand that their sum is the energy input flux required at the coronal base to maintain the loop stable. This should be a very short section, as it is a summary of one previous paper alone.

Due the difference between temperatures of transition region and solar corona, there must to be a heating mechanism that compensate all the losses present in the atmosphere and maintain stable the plasma temperature. Previous works have studied this mechanism by observations and modeling, but mostly for active regions. Mac Cormack *et al.* (2017) have developed a technique that estimate the heating injected into the loops during a period of minima activity. Using DGMT combined with an extrapolation of magnetic field, as we detailed before, they obtain the different plasma parameters along each magnetic loop, and therefore, compute the energy input flux required to maintain them stable. The model assume a simple energy balance for each magnetic flux tube, where the losses by radiative power (E_r) and thermal conduction power (E_c) are compensated by some coronal heating power (E_h) (Aschwanden, 2004):

$$E_h(s) = E_r(s) + E_c(s), \quad (7)$$

where s is the position along each magnetic flux tube and the powers are in units of $[\text{erg sec}^{-1} \text{cm}^{-3}]$.

NOTE: Trabajé todo lo naranja. Primero decidí que es mejor no numerar E_c sino F_c . Eso te quedará claro cuando agregues las ecuaciones (11) y (12) que te indico más abajo y leas el último párrafo naranja de abajo, ahí cierra todo. Además describí que son los ϕ en forma correcta. El texto previo estaba muy flojón, entre otras cosas se hablaba de una (inexistente) “conductive loss function” y de una (incorrectamente denominada) “area along the loop”. Obviamente luego de chequear este texto naranja borré esta nota azul en particular.

The thermal conduction power E_c equals the divergence of the conductive heat flux F_c , i.e. $E_c(s) = [1/A(s)] d[A(s) F_c(s)]/ds$, where $A(s)$ is the transversal area of the magnetic flux tube at position s . Under a quiescent solar corona plasma regime, the conductive flux is assumed to be dominated by the electron thermal conduction, described by the usual Spitzer model (Spitzer, 1962)

$$F_c(s) = -\kappa_0 T(s)^{5/2} \frac{dT}{ds}(s), \quad (8)$$

where $\kappa_0 = 9.2 \times 10^{-7} \text{erg sec}^{-1} \text{K}^{-7/2}$ is the Spitzer thermal conductivity.

Radiative power depends on the amount of plasma in certain temperatures that can radiate. The model estimate it by integrating the squared density multiplied by a radiative loss function $\Lambda(T)$. This function depends on plasma temperature and is calculated by the atomic database and the plasma emission model from CHIANTI (Del Zanna *et al.*, 2015). Then, the expression for radiative power obtained is:

$$E_r = \int dT \text{LDEM}(T) \Lambda(T) \quad (9)$$

The energy balance given by Equation (7) is then integrated in the volume of any given coronal magnetic flux tube. Dividing the result by the flux tube area at the coronal base, and making use of the soleidonal conditon of the magnetic field, a field-line integrated version of that energy balance is found,

$$\phi_h = \phi_r + \phi_c. \quad (10)$$

where the line integrated flux quantities $\phi_{r,c} [\text{erg sec}^{-1} \text{cm}^{-2}]$ are given by (Mac Cormack *et al.*, 2017),

$$\phi_r = \left(\frac{B_0 B_L}{B_0 + B_L} \right) \int_0^L ds \frac{E_r(s)}{B(s)} \quad (11)$$

$$\phi_c = \left(\frac{B_0 F_{c,L} - B_L F_{c,0}}{B_0 + B_L} \right) \quad (12)$$

Note that, for any given field line, all quantities in these two expresions can be computed from the DMT and AWSOM models through Equations (8)-(9). Once computed, the quantity ϕ_h can be calculated, which is the energy input flux required at the coronal base of each coronal field-line to maintain a stable coronal structure.

3. Results

3.1. Tomographic Results

NOTE: This section will show first the DMT products in the form of Carrington maps at three heights, for the two target rotations. It will then show the latitude/longitude location of the four types of magnetic loops (defined in Section 2.3) at a sample height. Then, for the each type of loop, we will show histograms of the values of different DMT products traced along them, characterizing in a quantitative fashion the DMT results in distinct magnetic structures. It will also include a quantitative comparison between the two target rotations of the median values of the DMT products in each type of loop.

As described in Section 2.1, we carried out DMT reconstructions of the coronal structure for target rotations CR-2082 and CR-2208 using STEREO/EUVI and SDO/AIA data, respectively. Once the LDEM was determined for each rotation, the square root of the mean value of the electron density squared $\sqrt{N_m^2}$ and the electron mean temperature T_m were computed at each voxel of the tomographic computational grid by means of Equations 2 and 3, and the measure R was calculated by means of Equation (4).

As an example, Figures 1 and 2 show latitude-longitude maps of DMT results for both rotations. Three different heights of interest are selected from the tomographic grid, providing also a detailed 3D view of the tomographic results: the lowest height of tomographic results ($1.025 R_\odot$), the lowest height where the AWSoM results are fully consistent with coronal conditions ($1.065 R_\odot$), and a middle height of the tomographic grid ($1.105 R_\odot$). Black voxels correspond to non-reconstructed voxels (see Section 2.1). Thick-black curves indicate the open/closed boundaries of the magnetic field of the AWSoM model.

Both target rotations are highly axisymmetric, i.e. characterized by a high azimuthal symmetry. Rotation CR-2082 showed two small ARs (active regions), both near latitude $+30^\circ$ and around longitudes 50° and 120° , respectively. Rotation CR-2208 showed two ARs, both near latitude $+5^\circ$ and around longitudes 140° and 300° , respectively.

The magnetically open and closed regions of the AWSoM model are associated to CHs and the equatorial streamer belt, respectively. Consistently with that, the location of the open/closed boundaries derived from the respective AWSoM model quite accurately match the regions of the DMT maps which exhibit the strongest latitudinal gradient of both the electron density and temperature.

Figures 1 and 2 show that the DMT reconstruction of the streamer belt is characterized by relatively larger values of density and temperature in comparison to the CHs. They also show that the streamer belt region of CR-2082 was denser and colder than that of CR-2208. In the case of CR-2082, which belongs to the deep minimum epoch between SCs 24 and 25, the low latitudes of the streamer belt are characterized by lower electron temperature than in its mid-latitudes. A similar behaviour is seen in CR-2208, belonging to the declining phase of SC 25, but having a somewhat less axisymmetric structure this characteristic is not so obvious. This thermodynamic structure of the streamer have been reported for other solar minimum rotations in previous DMT works

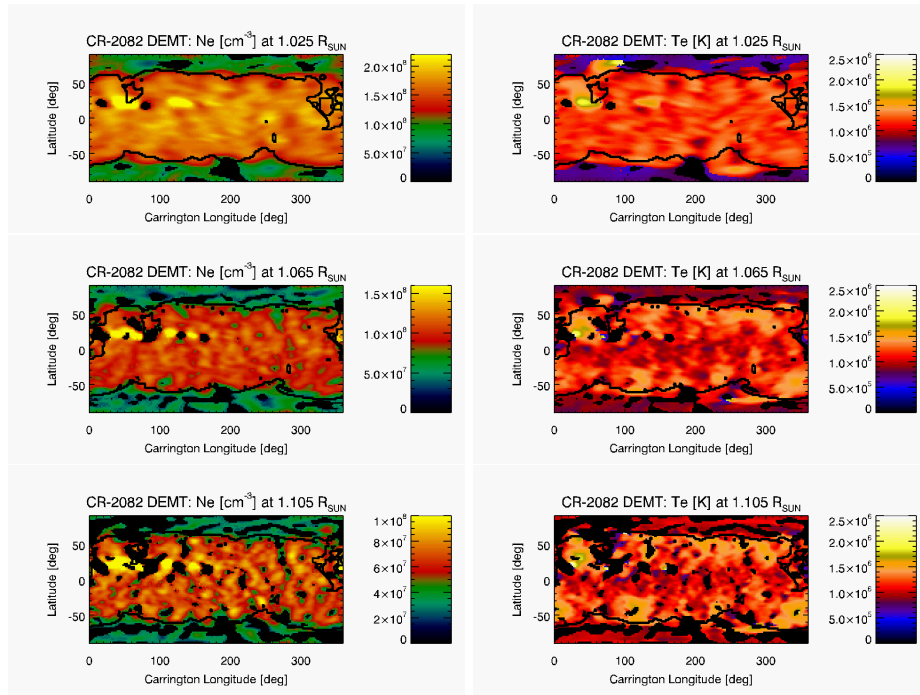


Figure 1. Carrington maps of DENT products $\sqrt{N_m^2}$ (left panels) and T_m (right panels) for CR-2082. Top, middle and bottom panels show the results at three heliocentric heights, 1.025, 1.065 and 1.105 R_\odot respectively. Black voxels correspond to non-reconstructed regions (see text in Section 3.1) and thick-black curves indicate the open/closed boundaries.

(Lloveras *et al.*, 2017; Nuevo *et al.*, 2013; Vásquez, Frazin, and Manchester, 2010).

For both target rotations, Figure 3 shows latitude-longitude maps of the DENT measure R defined by Equation (4), at the same three heights shown in Figures 2 and 3. In most of the coronal volume of the DENT grid the agreement between the tomographic and synthetic FBEs is $\lesssim 1\%$. The notable exception is to be found in the CHs of the target rotation analysed based on AIA data. A similar results was found in the two existing DENT works based on data provided by the AIA instrument (Nuevo *et al.*, 2015; Mac Cormack *et al.*, 2017). This point is further discussed below.

To characterize the DENT results in different magnetic structures, we traced $\sqrt{N_m^2}$ and T_m along the magnetic field lines of the AWSoM model. For both rotations, all field lines that meet the criteria listed in Section 2.3 were selected. For each field line, the data points of electron density and electron mean temperature as a function of height were fitted to the Equations 5 and 6. As a result, the electron density $N_0(r = 1.0 R_\odot)$ and scale height λ_N were computed for each leg, as well as the temperature gradient $\nabla_r T_m$, and the height-averaged (along the leg) electron temperature $\langle T_m \rangle$.

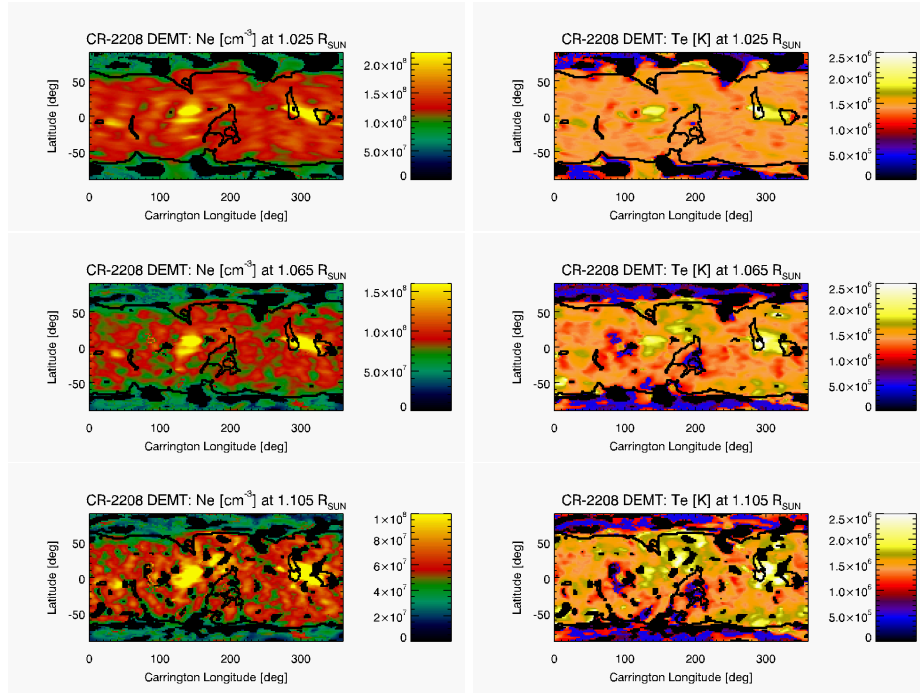


Figure 2. Same as Figure 1 but for CR-2208.

For both target rotations, the top panels of Figure 4 show the latitude-longitude location (at heliocentric height $r = 1.105 R_{\odot}$) of all traced field line legs for which criterion (i) of Section 2.3 is met. Open legs are indicated in grey color and closed ones in black color. For each leg, the fits to tomographic temperature and density were applied, as given by Equations (5)-(6). Considering the DEMT data points and the resulting fits along each leg, the bottom panels of Figure 4 show the latitude-longitude location of the subset for which also both criteria (ii) and (iii) of Section 2.3 are met. Using a four-color code, type 0, I, II and III legs are shown in blue, red, magenta and cyan color, respectively. Of the 45000 legs selected for CR-2082, 14% are type 0, 38% are type I, 23% are type II and 25% type III. On the other hand, of the 59000 legs selected for CR-2208, 8% are type 0, 43% are type I, 22% are type II and 30% type III. **NOTE: Los números deben ser actualizados, o la frase eliminada.**

Type 0 (small down) legs mainly populate the equatorial latitudes. This kind of structure was originally found by Huang *et al.* (2012), and their existence was shown to be anti-correlated with the solar activity level around the solar minimum between SCs 24 and 25 by Nuevo *et al.* (2013). Later on, Lloveras *et al.* (2017) showed that equatorial down loops in streamers were also to be found in the deep minimum between SCs 23 and 24. Here, we verify the existence of this type of structure for the two target rotations. The relatively smaller population of down loops seen in CR-2208, as compared to CR-2082, is consistent with the aforementioned results by Nuevo *et al.* (2013). Type I (small up) mainly

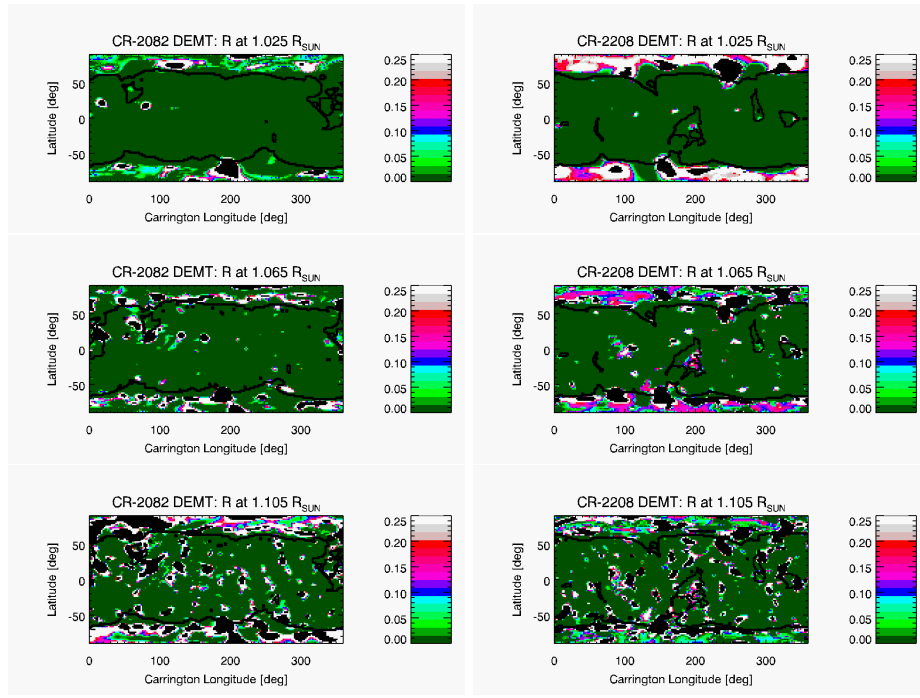


Figure 3. Carrington maps of the measure R defined by Equation (4), for CR-2082 (left panels) and CR-2208 (right panels), at heights 1.025, 1.065 and $1.105 R_{\odot}$, from top to bottom.

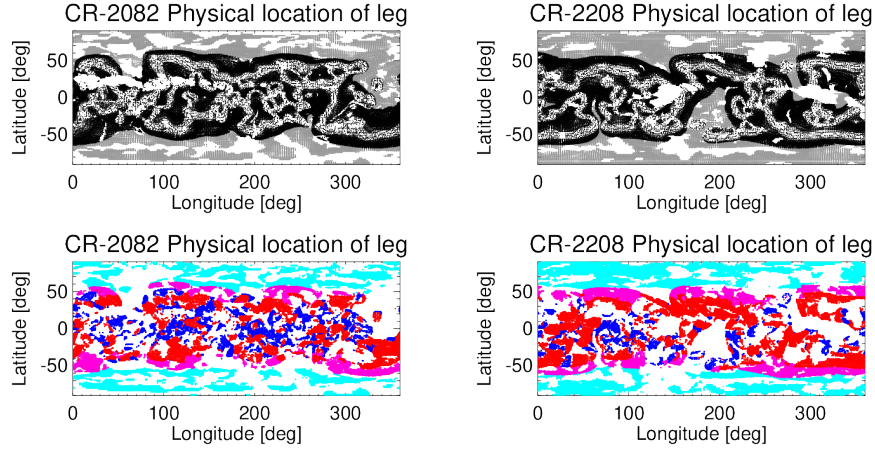


Figure 4. *Top panels:* latitude-longitude location at heliocentric height $r = 1.105 R_{\odot}$ of all open (grey color) and closed (black color) traced field line legs for which criterion (i) of Section 2.3 is met, for both CR2082 (left) and CR2208 (right). *Bottom panels:* latitude-longitude location of the subset for which also both criteria (ii) and (iii) of Section 2.3 are met. The location of type 0, I, II and III legs is shown in blue, red, magenta and cyan color, respectively.

populate the mid-latitudes, while type II (large up) legs are mostly very large trans-equatorial field lines forming the envelope of the streamer belt. Finally, type III (open) legs populate of course the CHs.

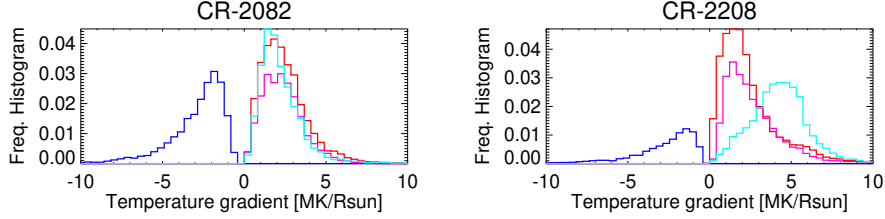


Figure 5. Frequency histograms of the temperature radial gradient for the four types of legs in Figure 4 (using the same color code) for CR-2082 (left panel) and CR-2208 (right panel).

Figure 5 shows frequency histograms (normalized to the total amount of selected field lines) of the temperature radial gradient ($\nabla_r T_m$) for legs of type 0, I, II and III. The lack of population around values close to zero is due to the requirement $|\rho_t| > 0.5$ which discards quasi-isothermal legs. For both rotations, the median value of the temperature radial gradient is $\text{Md}(\nabla_r T_m) \approx -2.5, +2.3$ and $+2.4 \text{ MK/R}_\odot$ for legs of type 0, I and II, respectively.

The notable difference between both rotations is the characteristic value $\text{Md}(\nabla_r T_m) \approx +4.5 \text{ MK/R}_\odot$ for legs of type III for CR-2208. This is related to the much larger R score for the DENT results along CH legs for CR-2208. In this case, DENT performs poorly in modeling a LDEM that predicts the tomographic FBEs with reasonable accuracy. Indeed, visual inspection of the DENT temperature maps for CR-2208 in Figure 3, reveals that in most of the CH region the result for T_m below height $1.105 R_\odot$ is quite uniform and artificially low. As it turns out, around and above height this height the score R is lower and the temperature results are more reliable. As a result, the temperature gradient along these legs is artificially larger, with the linear fit trying to simultaneously fit the artificially low values of T_m at lower heights. In general, the DENT results in the CH region based on AIA data are thus much less reliable than in the rest of the analysis. We will return to this point in the conclusion section.

For both rotations, Figure 6 shows, in a statistical fashion, the DENT results traced along field lines discriminated by field line type. From top to bottom results are shown for type 0 to type III field lines, respectively. From left to right the panels show the statistical distribution of $N_{CB} \equiv \sqrt{N_m^2}(r = 1.055)$ (the lowest height where the AWSOM results are consistent with coronal conditions), λ_N and $\langle T_m \rangle$, with the median value m indicated in each plot.

Table 1 summarizes a quantitative comparative analysis between the results of the two target rotations. For CR-2082 quantities are expressed as absolute values, while for CR-2208 they are informed as a precentual variation relative to the corresponding results for CR-2082. The following major results, both concerning the structure of each rotation individually as well as their comparison, can be drawn.

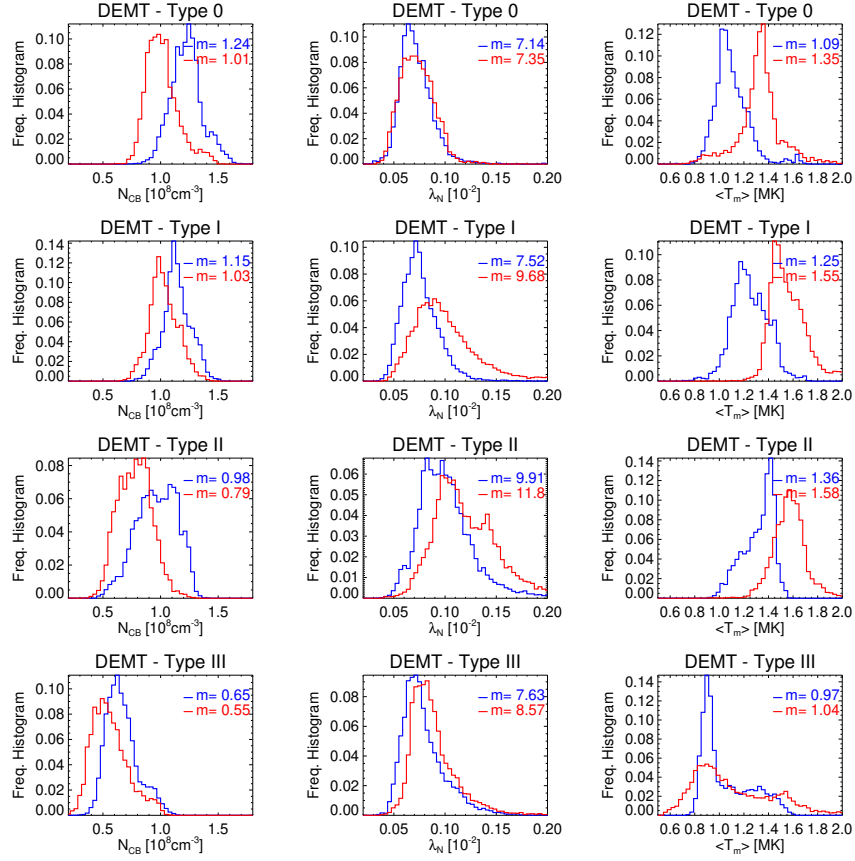


Figure 6. Statistical distribution of DENT results for rotations CR-2082 (blue) and CR-2208 (red) traced along legs of type 0, I, II and III (from top to bottom), as defined in Section 2.3. From left to right: electron density $N_{CB} \equiv \sqrt{N_m^2}(r = 1.055 R_\odot)$, electron density scale height λ_N , and loop-averaged temperature (T_m). In each panel the median value m is indicated.

Throughout the magnetically closed region of both rotations, type 0, I and II legs, associated to increasingly outer regions of the equatorial streamer belt, progressively exhibit decreasing coronal base density, increasing density scale height, and increasing electron temperature. In both rotations also, type III field lines in the CHs are characterized by sub-MK temperatures, and electron density values of order $\approx 1/2$ of those observed for the type 0 and type I lines in the core of the equatorial streamer.

A comparison of the results between the two rotations shows that, compared to CR-2082, target rotation CR-2208 was characterized by $\approx 15 - 20\%$ lower values of the electron density at the coronal base, $\approx 10 - 30\%$ larger values of density scale height, and $\approx 15 - 25\%$ larger values of the electron temperature. **NOTE:** Los números de esta frase deben ser consistentes con los que finalmente informe la Table 1, redondeando los rangos característicos observados para los

Table 1. Median value (indicated as “Md”) of the statistical distribution of N_{CB} , λ_N , and $\langle T_m \rangle$ for each coronal type of lines defined in Section 2.3. For CR-2082 values are expressed in absolute terms, while for CR-2208 they are informed as a percentual variation relative to the CR-2082 value. NOTE: Propongo que para todas las cantidades absolutas de la Tabla 1 informes siempre dos cifras significativas, y que todos los porcentajes los redondees a cero decimales (o sea, a veces una y a veces dos cifras significativas) y que no uses puntos para estos nunca. Doy dos ejemplos para no dejar dudas. Para N_{CB} de Type-II: 1.0 (−21%). Para $\langle T_m \rangle$ de Type-III: 1.0 (+6%).

Type	Md(N_{CB}) [10^8 cm^{-3}]	Md(λ_N) [$10^{-2} \text{ R}_{\odot}$]	Md($\langle T_m \rangle$) [MK]
0	1.26 (−17.4%)	7.3 (+13.6%)	1.09 (+24.7%)
I	1.13 (−14.1%)	8.3 (+18.1%)	1.29 (+20.1%)
II	0.97 (−20.6%)	8.9 (+29.1%)	1.33 (+16.5%)
III	0.66 (−16.2%)	7.7 (+11.4%)	0.98 (+ 6.1%)

diversos tipos de línea (consultame!). La Table 1 debe a su vez resumir lo que muestra la Figure 6. Destacalo en tu color cuando lo hayas hecho así le presto atención.

To analyze the loop-integrated energy flux quantities introduced in Section 2.3, we selected closed loops for which both legs have the same sign of the radial gradient of the electron temperature $\nabla_r T_m$. In this way, according to the clasification of both its legs, each given loop was classified as of type 0 (small down loop), I (small up loop), or II (large up loop). For both target rotations, and for loops of type 0, I and II, Figure 7 shows the frequency histogram of the loop-integrated energy flux quantities ϕ_r , ϕ_c and ϕ_h in blue, green and red color, respectively.

For both rotations, the value of the **loop-integrated** radiative power E_r , measured by the quantity ϕ_r , is largest for loops of type 0. This is due to $E_r \propto N_e^2 \Lambda(T_e)$, with both factors contributing to maximize E_r for loops of type 0. As shown in Figure 6 and Table 1, loops of type 0 are characterised by **the largest** values of electron density. Also, in the range of sensitivity of the EUVI and AIA instruments, namely 0.5–3.0 MK (Nuevo *et al.*, 2015), the radiative loss function $\Lambda(T)$ has a local maximum at $T_c \approx 1$ MK. According to Figure 6, loops of type 0, I and II are characterised by values of temperature that are progressively larger and farther from the value T_c , for both rotations.

The sign of the quantity ϕ_c depends on that of the conductive flux F_c . Equations (8) and (12) imply that down loops (type 0) and up loops (type I and II) are characterized by $\phi_c < 0$ and $\phi_c > 0$, respectively, **as verified** in Figure 7.

Adding the radiative and conductive terms, the characteristic energy input flux at the coronal base is in the range $\phi_h \approx 0.5 - 1.5 \times 10^5 \text{ erg cm}^{-2} \text{ s}^{-1}$, depending on the rotation and the type of **loop, matching the values reported by** Mac Cormack *et al.* (2017). Note that for type 0 loops there is a marginal population characterized by the unphysical result $\phi_h < 0$. As shown by Mac

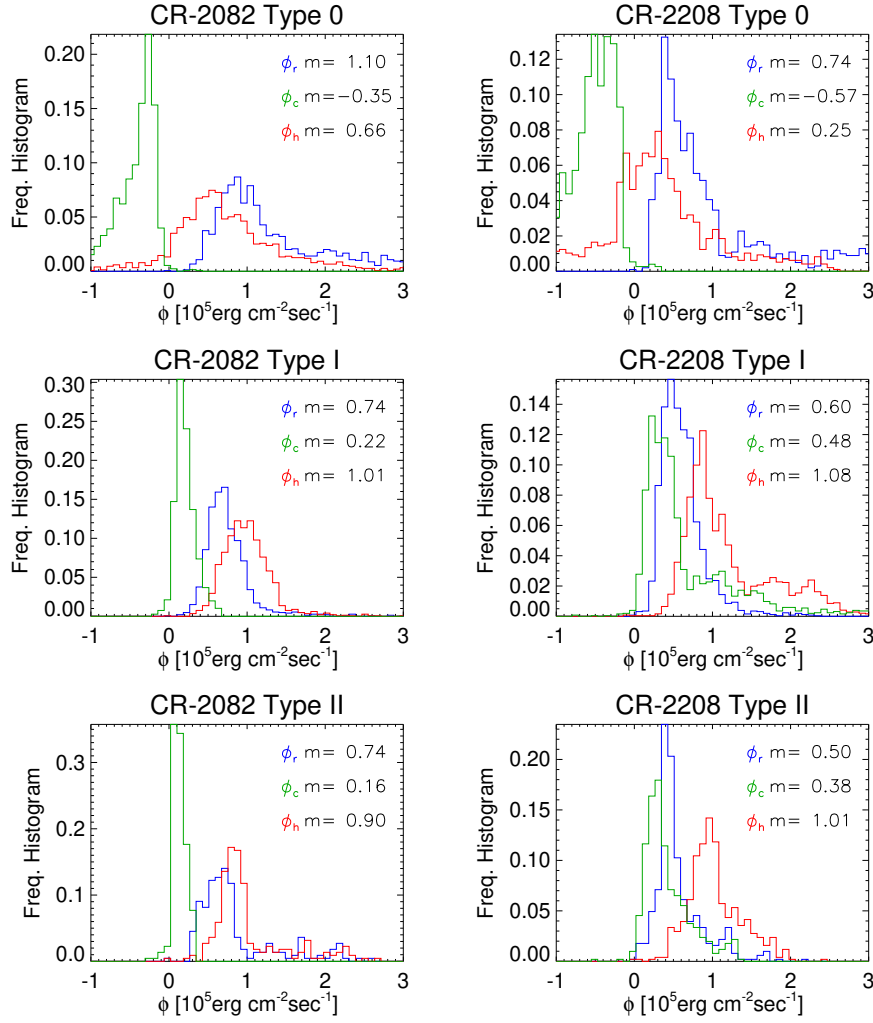


Figure 7. Statistical results of the loop-integrated energy flux quantities ϕ_r , ϕ_c , and ϕ_h in colors blue, red and green, respectively for CR-2082 (left) and CR-2208 (right). From top to bottom, panels show the results for loops of type 0, I and II, which are loops for which both legs meet the criteria from Section 2.3.

Cormack *et al.* (2017), this affects only the smallest sized loops of the type 0, and it is most probably due to the limited temperature sensitivity of the instrumental passbands. The radiative loss term is here calculated based on plasma emission detected by three coronal bands of EUVI or AIA. Though accounting for most of the coronal plasma, there surely is additional emission out of the instrumental sensitivity range. As a result, the positive term ϕ_r is most likely underestimated, leading to values $\phi_h < 0$ in loops of type 0, being characterized by $\phi_c < 0$.

3.2. Tomography and Models

For both target rotations, Figures 8 and 9 show latitude-longitude maps of the AWSoM electron density and temperature. Maps are shown at the same three heights selected for visualisation of the DMT results in Figures 1 and 2. Thick-black curves indicate the magnetic open/closed boundaries based on the magnetic field of the AWSoM model. Visual inspection of these maps shows that the AWSoM model for both rotations is highly axisymmetric, as the tomographic model.

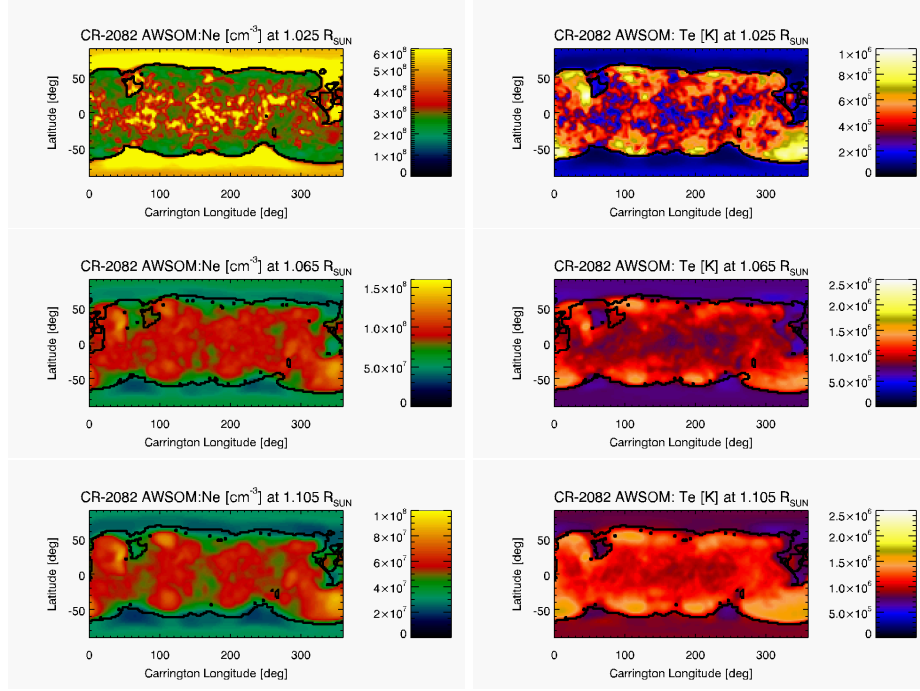


Figure 8. Carrington maps of density (left panels) and temperature (right panels) obtained with AWSoM model for the same three heights shown in Figure 1.

As described in Section 2.2, the AWSoM model includes an artificially thick TR, achieving coronal conditions above height $\approx 1.06 R_{\odot}$. Indeed, the top panels in Figures 8 and 9, at height $1.025 R_{\odot}$, clearly do not represent coronal conditions (although we include them here for completeness). When compared to DMT results (Figures 1 and 2), the latitude-longitude maps of the AWSoM model for heights 1.065 and $1.105 R_{\odot}$ capture well the denser and hotter equatorial streamer belt surrounded by the less dense and colder CHs. Furthermore, for both target rotations, the temperature maps show the low latitudes of the equatorial streamer belt to be characterised by lower temperatures than its mid-latitudes, as also seen in the DMT results. The latitude-longitude maps of the AWSoM and DMT results are shown in the same units and scales, so that a

visual comparison among them already reveals similar values of electron density and temperature in both models.

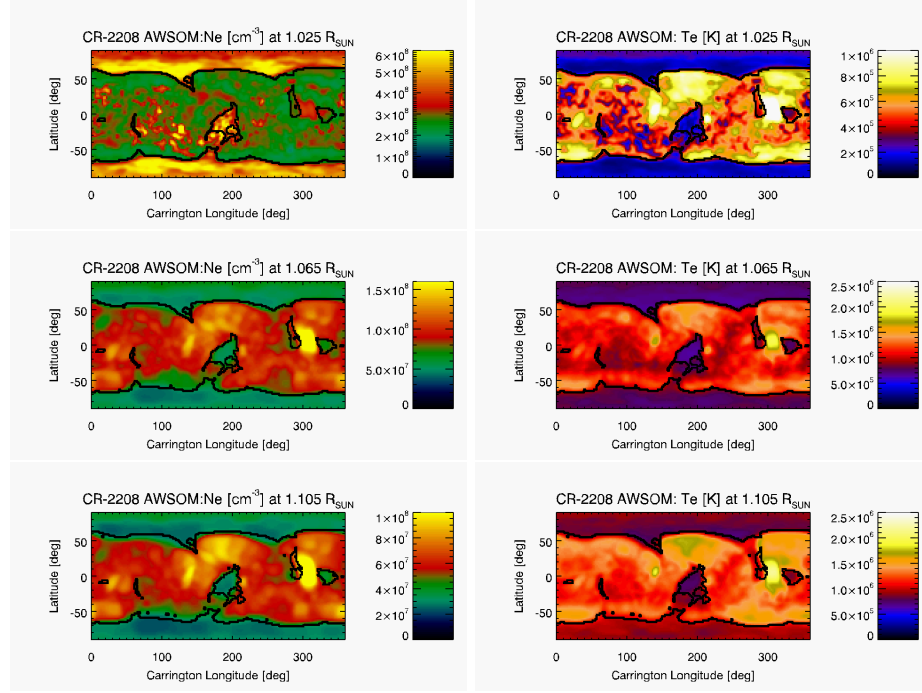


Figure 9. Same as Figure 9 for CR-2208.

Being highly axisymmetric target rotations, the longitude-averaged latitudinal variation of the results of both models is an informative way to compare their large-scale structure. Such comparison is shown in Figure 10 at height 1.105 R_{\odot} , with top panels comparing electron density and lower panels electron temperature. In these longitude-averaged profiles, longitudes containing ARs or low latitude CHs were excluded. In each panel the averaged latitudinal variation for the DMT model is shown in solid-line style, while the result for the AWSOM model is shown in dashed-line style. Left panels show the comparison for CR-2082 (in blue) and right panels for CR-2208 (in red). In each panel the vertical black lines denote the corresponding longitude-averaged latitude of the open-close boundary in both hemispheres.

NOTE: Diego: debes estar seguro que sacaste ARs y low-latitude CHs de los promedios en longitud. Cuando más te esmeres la comparación es más justa, y podrás revelar algo que de otra manera no se vería. Si promediaste longitudes con ARs o CHs, rehacelo, vale la pena. Por supuesto debes promediar idénticas longitudes en DMT, AWSOM y en la O/C.

Several details from Figure 10 are worth being highlighted. Firstly, for CR-2082 the electron density of both models agree within $\approx 20\%$ at all latitudes, and for CR-2208 the agreement is within $\approx 5\%$. In the case of the electron

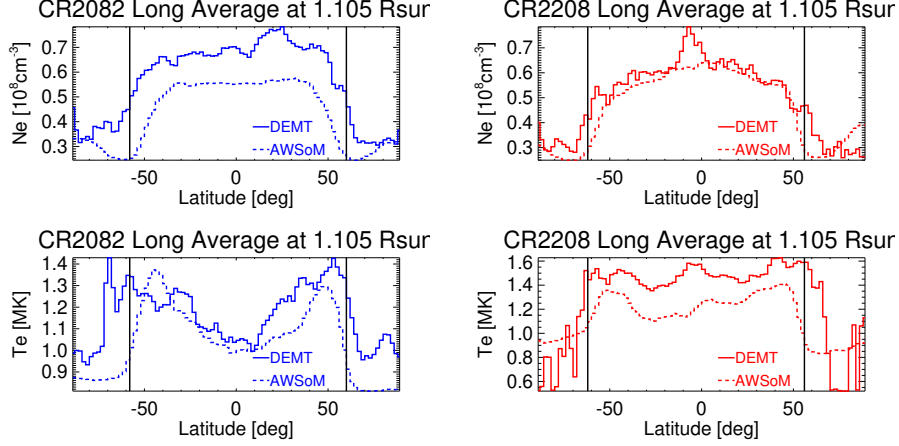


Figure 10. Longitude-averaged latitudinal dependence of the electron density (top) and temperature (bottom) for DMT (blue) and AWSOM (red) results at $1.105 R_{\odot}$. The left (right) panels correspond to CR-2082 (CR2208). The vertical black line indicates the longitude-averaged latitudes of the open/close magnetic boundary in both hemispheres.

temperature both models agree within $\approx 10\text{--}15\%$ at all latitudes for both target rotations. Secondly, for both target rotations, and for both models, these plots clearly show the relatively lower temperatures characterising the low-latitudes of the equatorial streamer belt compared to its mid-latitudes. Thirdly, for both target rotations, the latitude of the open/close magnetic boundary in both hemispheres matches the location of the strongest latitudinal gradient of the DMT electron density. Note this is not the case for the AWSOM model, that shows a minimum density at the open/closed boundary. Last, note that the DMT electron density decreases from the open/close boundary towards the poles (in both hemispheres of the two target rotations), while the AWSOM model shows the opposite trend.

To characterize the results of the AWSOM model in distinct magnetic structures, its results for electron density and temperature were traced along its magnetic field lines. For each field line leg, the results were then fit to Equations (5) and (6). In its current implementation, the AWSOM model can not simulate down loops, a point we will discuss in the next section. We then classified the traced legs into types I, II and III, according to the criteria described in Section 2.3.

For both target rotations, the top panels of Figure 11 show the latitude-longitude location (at heliocentric height $1.105 R_{\odot}$) of all traced field line legs for which criterion (i) of Section 2.3 is met. That criterion is adapted here, requiring that at least five voxels of the tomographic grid are threaded by the leg. Open legs are indicated in grey color and closed ones in black color. For each leg, the fits to tomographic temperature and density were applied, as given by Equations 5 and 6. Considering the AWSOM data points and the resulting fits along each leg, the bottom panels of Figure 11 show the latitude-longitude location of the

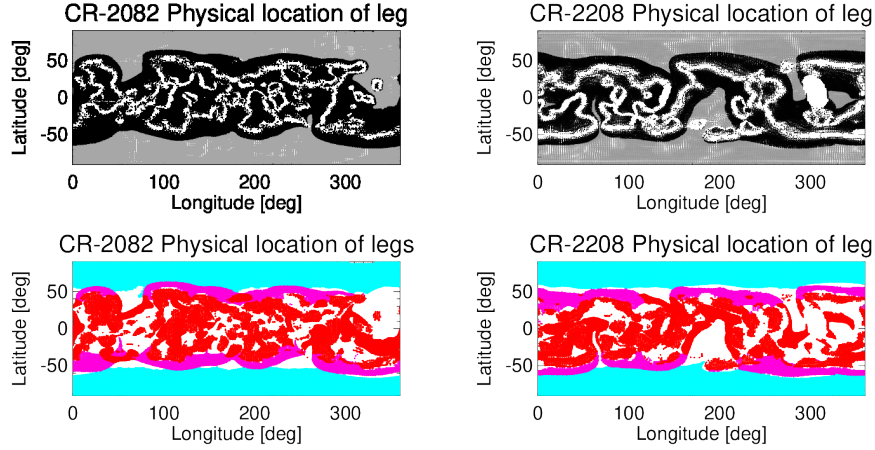


Figure 11. Same as Figure 4, but using the density and temperature of the AWSoM model to classify its legs in types I, II and III. The model does not exhibit legs of type 0.

subset for which also both criteria (ii) and (iii) of Section 2.3 are met. Using a three-color code, type I, II and III legs are shown in red, magenta and cyan color, respectively. This figure is to be compared with the corresponding Figure 4 for DGMT results. It is readily seen that the AWSoM maps are more populated than those of DGMT. This is due to the 3D MHD model having electron density and temperature fields spatially smoother than those of DGMT.

For target rotation CR-2082, Figure 12 shows the statistical distribution of the results of the DGMT (solid line-style) and AWSoM (dashed line-style) models traced along legs of type I, II and III (from top to bottom), as defined in Section 2.3. From left to right: electron density at the lowest coronal height of the AWSoM model $N_{CB} \equiv N_e(r = 1.055 R_\odot)$, electron density scale height λ_N , and leg-averaged electron temperature $\langle T_m \rangle$. In each panel the median values m are indicated. Figure 13 shows the same analysis for target rotation CR-2208.

For the two target rotations, Table 2 summarizes a quantitative comparative analysis between the results of the DGMT and AWSoM models. The DGMT results are expressed as absolute values, while the AWSoM results are informed as a percentual variation relative to the corresponding result for DGMT. the following major results can be drawn.

For target rotation CR-2082, the median value of the electron density N_{CB} of both models agree within $\approx 10 - 25\%$, depending of the type of leg, while the median value of the scale height λ_N agree within $\approx 10\%$. The leg-averaged electron temperature $\langle T_m \rangle$ of both models also agree within 10%. For target rotation CR-2208 the agreement of the median value of N_{CB} and λ_N of both models is within 5%, while median values of $\langle T_m \rangle$ agree within 15%. These detailed results informed per magnetic structure are fully consistent with the large-scale comparison provided in Figure 10.

Finally, to provide a graphical comparison of both models accross the full range of heliocentric heights covered by the DGMT results, Figure 14 shows the

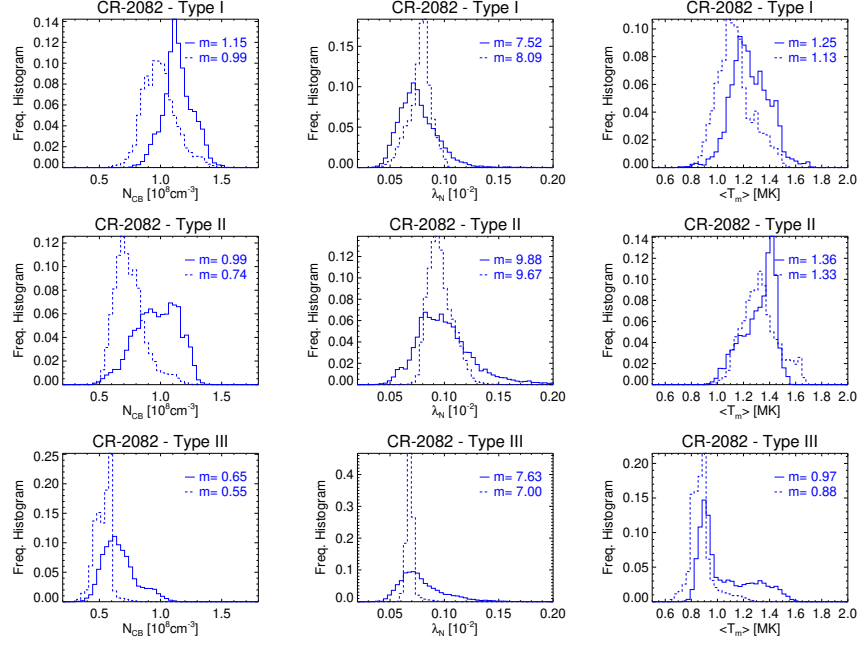


Figure 12. Statistical distribution of the results of the DEMT (solid line-style) and AWSOM (dashed-line-style) models traced along legs of type I, II and III (from top to bottom), as defined in Section 2.3. From left to right: electron density at the lowest coronal height of the AWSOM model $N_e(r = 1.055 R_\odot)$, electron density scale height λ_N , and leg-averaged electron temperature $\langle T_m \rangle$. In each panel the median values m are indicated.

average fits of $N_e(r)$ and $T_e(r)$ for legs of type I (red), II (magenta), and III (cyan) for both target rotations. In each panel the DEMT and AWSOM results are plotted in solid and dashed linestyles, respectively.

As discussed above, Figure 10 shows that the longitude-averaged latitudinal profile of the DEMT electron density in the CHs decreases towards the poles. Figure 15 below shows the longitude-averaged AWSOM radial wind speed V_r at $6 R_\odot$, where all field lines are open. The heliocentric current sheet (HCS) location is indicated by the minimum of the speed curve. For each rotation, all velocity data points to the south of the HCS position map down to the southern CH in Figures 10. Similarly, all velocity data points to the north of the HCS position map down to the northern CH in Figures 10. This clearly shows an anti-correlation between the DEMT electron density at low heights and the AWSOM wind speed at larger heights.

NOTE: We could quantify the anti-correlation between the AWSOM terminal speed and the DEMT electron density at the coronal base in the following fashion. We can set up starting points at, say, $20 R_\odot$, every one degree, both in latitude and longitude. This would mean setting up 180×360 strating points. We trace then the (open) field lines from each of those starting points down to $1.0 R_\odot$ to find out the DEMT results at those lower heights for each field line. Now, to do this we need to develop some codes and have no time for this paper.

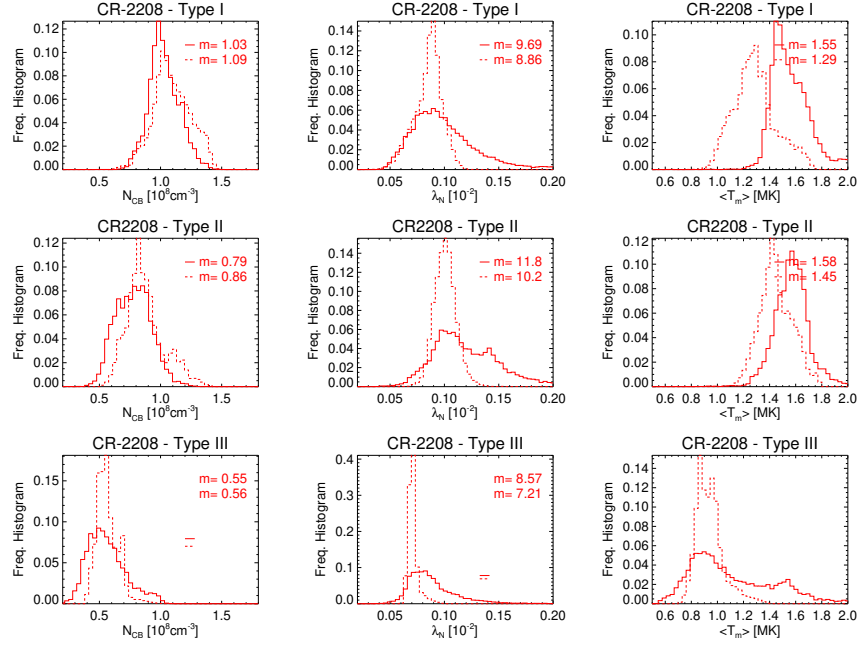


Figure 13. Same as Figure 12 for CR-2208.

We can postpone it for a next effort, or... Chip: do you think you could do this with your present tools? Do you think it is interesting? I think so. Our submission deadline is end-of-January.

4. Discussion

NOTE: (by Albert) This section will discuss and summarize some main conclusions, possibly in three groups: a) DENT analysis alone (including thermodynamical structure, up/down loops, energy considerations, etc.), b) Comparison ASoM/DENT (longitude-averaged latitudinal profiles, height-dependence profiles in different magnetic structures, etc), and c) Correlation between ASoM terminal wind speed and DENT basal density, I will expand on this later this week as this may require help from Michigan).

NOTE: Here is a paragraph describing the empirical successes in the reconstruction of both rotations.

*In comparing the DENT results obtained for the two selected targets, it is important to bear in mind they rely on data provided by two different instruments: AIA and EUVI. In order to quantify the systematic difference of the DENT products based on both instruments, Nuevo *et al.* (2015), who were the first to apply DENT to AIA data, analysed a single target using both instruments independently. They found that while the density product is essentially

Table 2. Median value (indicated as “Md”) of the statistical distribution of N_{CB} , λ_N , and $\langle T_m \rangle$ for each coronal type of lines defined in Section 2.3. DENT values are expressed in absolute terms, while AWSOM results are informed as a percentual variation relative to the corresponding DENT value. Propongo que para todas las cantidades absolutas informes siempre dos cifras significativas, y que todos los porcentajes los redondees a cero decimales (o sea, a veces una y a veces dos cifras significativas) y que no uses puntos para estos nunca. Ejemplo para CR-2082 N_{CB} de Type-II: 1.0 (−23%). Ejemplo para CR-2208 λ_N de Type-II: 12.0 (−17%). Ejemplo para CR-2208 $\langle T_m \rangle$ de Type-II: 1.6 (−10%).

Type	Md(N_{CB}) [10^8 cm^{-3}]	Md(λ_N) [10^{-2} R_\odot]	Md($\langle T_m \rangle$) [MK]
CR-2082			
I	1.13 (−11.9%)	8.3 (− 2.4%)	1.29 (−10.8%)
II	0.97 (−22.5%)	8.9 (+ 2.2%)	1.33 (− 4.5%)
III	0.66 (−16.8%)	7.7 (− 9.1%)	0.98 (− 9.2%)
CR-2208			
I	0.97 (+ 9.3%)	9.8 (− 7.1%)	1.55 (−14.2%)
II	0.77 (+11.7%)	11.6 (−17.2%)	1.55 (− 9.7%)
III	0.56 (− 1.9%)	8.6 (−16.3%)	1.04 (− 9.6%)

equal, the temperature product of DENT based on AIA data is systematically 8% larger than the one based on EUVI data, i.e. $T_m^{(AIA)}/T_m^{(EUVI)} \approx 1.08$. Considering such correction, Figure 6 and Table 1 indicate that CR-2208 was $\approx 10 - 15\%$ hotter than CR-2082 throughout the streamer belt region. As for the electron density products, CR-2208 was found to be $\approx 15 - 20\%$ less dense than CR-2082 throughout the streamer belt region. These systematic differences are beyond the uncertainty level in the DENT products due to systematic sources (radiometric calibration and tomographic regularization), that Lloveras *et al.* (2017) estimated to be $\lesssim 5\%$ and $\approx 2\%$ for the electron temperature and density, respectively.

We found down legs in the DENT reconstructions mainly in the equatorial latitudes, being a greater quantity and better distribution along longitudes in CR-2082 than in CR-2208. Nuevo *et al.* (2013) showed a decrease in the amount of down loops as the solar activity increases, which is consistent with the results shown in this article.

*We can see, on type 0 loops, a median value of energy input flux (ϕ_h) for CR-2082 twice bigger than CR-2208. In both rotations this magnitude has negative values, this can refer to some loops where loop-integrated radiative flux (ϕ_r) is too small in comparison with the larger gradients of negative conductive flux (ϕ_c), resulting in negative energy input flux.

This effect can be due a subestimation of the density present in the loop or the limited range of temperature used to reconstruct tomographic parameters, that can affect the compute of radiative power. Nevertheless, this population

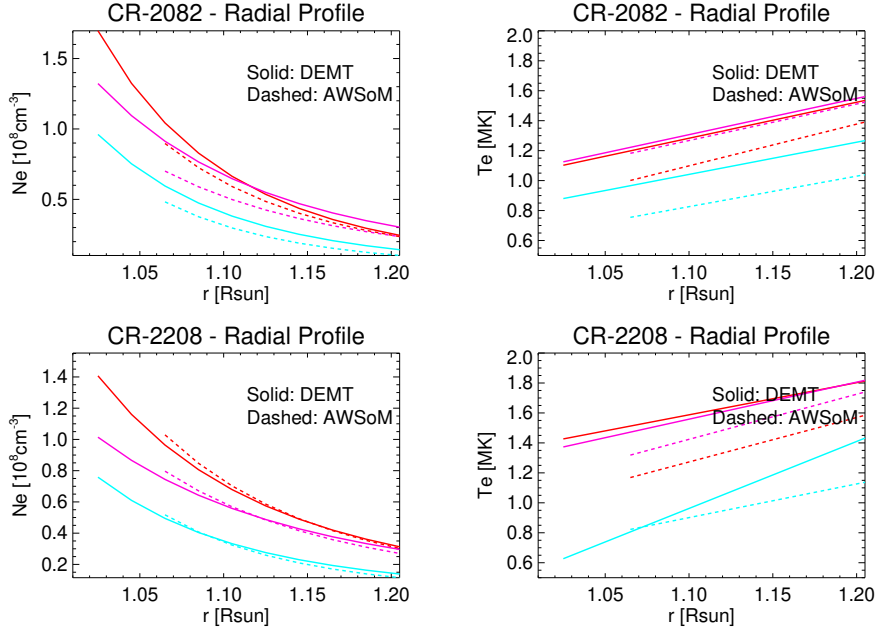


Figure 14. Average fits to $N_e(r)$ (left panels) and $T_e(r)$ (right panels) for legs of **type I** (red), **II** (magenta), and **III** (cyan), for CR-2082 (top panels) and CR-2208 (bottom panels). Solid lines correspond to DMT results while dashed lines correspond to AWSoM results.

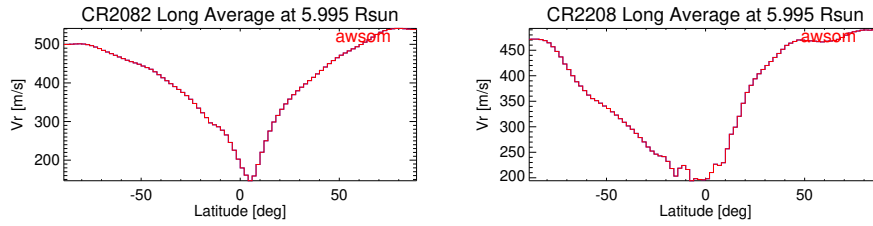


Figure 15. Longitude-averaged latitudinal dependence of the AWSoM model wind speed at $6.0 R_{\odot}$ for CR2082 (left panel) and CR2208 (right panel). **NOTE:** (by Albert): Even if this is not yet the terminal speed, it is a pretty high height, and one can already see the transition from FAST to SLOW speed in the model, where the absolute minimum indicates the location of the HCS at this height. I believe it is interesting to compare these two curves with Figure 10. Note that while in this Figure all latitudes are magnetically open, in Figure 10 only the part at latitudes larger than the vertical black lines are open. Note that DMT shows that in the open low corona N_e decreases from the O/C boundary towards the poles, anti-correlating with the AWSoM terminal speed, which is nice. In order to quantify this anti-correlation we need to trace the terminal speed and the DMT N_e , which we may need to do with the help of Chip and Nishtha. Chip, we can discuss this idea over Skype after Xmas if you are around.

represent lower than the **ALGÚN PORCENTAJE%** of total population in both cases.

Consistent that the density in CR-2082 was higher, the ϕ_r was also higher. In both rotations the magnitudes in the temperature gradients are similar, but the temperatures in CR-2208 were consistently higher at all heights, resulting in a larger ϕ_c in module. This results are in agree with the observations of Mac Cormack *et al.* (2017).

THE AR of CR-2082 wasn't well modeled by AWSoM.

In both rotations, AWSoM and DENT results shows same thermodynamic structures along latitude and longitude with very comparable values between $1.055 - 1.25 R_\odot$

Figure 12 and Table 2 indicate that AWSoM reconstruction of CR-2082 was found to be less dense $\approx 11 - 22\%$ and colder $\approx 5 - 10\%$.

Both reconstructions of CR-2208 reconstructed the ARs with very similar values and in almost the same latitude-longitude location. Figure 13 and Table 2 indicate that AWSoM reconstruction of CR-2208 was found to be denser $\approx 0 - 10\%$ and colder $\approx 10 - 15\%$.

CR-2208 was $\approx 10 - 15\%$ hotter than CR-2082 throughout the streamer belt region. As for the electron density products, CR-2208 was found to be $\approx 15 - 20\%$ less dense than CR-2082 throughout the streamer belt region. These systematic differences are beyond the uncertainty level in the DENT products due to systematic sources (radiometric calibration and tomographic regularization), that Lloveras *et al.* (2017) estimated to be $\lesssim 5\%$ and $\approx 2\%$ for the electron temperature and density, respectively.

I would like to say something about Figure 10. Maybe the weird tale in AWSoM Ne in the poles. Ideas?

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