

Thermodynamic Structure of the Solar Corona: Tomographic Reconstructions and MHD Modeling

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Abstract Observational techniques play an essential role in advancing our understanding of the physics of the solar corona. They provide validation data

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for three-dimensional (3D) magnetohydrodynamic (MHD) models of the solar atmosphere, key to improve their space weather forecasting capabilities. Solar rotational tomography (SRT) is currently the sole observational technique that provides an empirical 3D description of some of the fundamental plasma parameters of the solar corona at a global scale. Based on EUV data of space borne instruments, SRT allows constructing 3D maps of the coronal electron density and temperature at heliocentric heights below 1.25 R_{sun}. We carry out a study of the corona combining tomographic reconstructions with MHD simulations using the latest version of the Alfvén Wave Solar Model (AWSoM) of the Space Weather Modeling Framework (SWMF). Target rotations were selected from the solar minimum between solar cycles (SC) 23 and 24 and the current declining phase of solar cycle 24. Tomographic reconstructions and results of the model are analyzed in distinct coronal magnetic structures. We study magnetically closed structures associated with the equatorial streamer belt, as well as magnetically open regions enclosing it. We report on the tomographic results in the different structures, their implications for the physics of the solar corona, and the current capability of the AWSoM model to reproduce the tomographic reconstructions in different regions.

Keywords: Solar Cycle, Observations; Corona,E; Corona, Structures

1. Introduction

Being the place where the solar atmosphere is heated and the solar wind accelerated, and where impulsive events such as solar flares and coronal mass ejections are energized, observation and modeling of the solar corona are tasks of great relevance to improve our understanding of the Sun-Earth environment. To advance our knowledge of the physics of the solar corona, as well as to improve its three-dimensional (3D) models, information derived from observational data plays a key role. Solar rotational tomography is currently the sole observational technique able to provide a quantitative empirical description of the 3D distribution of some fundamental plasma parameters of the solar corona at a global scale.

NOTE: Here we will put in a couple of paragraphs summarizing what SRT and DENT are, including references and defining the accronyms in the process. We should refer to minima in particular. We should mention the instruments and missions we use as data for DENT, and define their accronyms too. Also, a couple of paragraphs on the AWSoM model and the SWMF, including references and defining the accronyms in the process.

2. Methodology

2.1. DENT reconstructions

NOTE: This section will summarize DENT's main aspects and equations, only those needed to understand the paper. Nothing not needed to read this paper is

provided. Also, no derivation of equations are provided. References are given so that the reader can dig deeper if so desires. Points to cover are:

- Selected targets, instruments used. One paragraph.
- Characteristics of the data (specific time series, dates, bands, etc.), SolarSoft software version used for their processing. One paragraph.
- Technical aspects of DENT (grid, regularization, off-limb data is only used) and how these differ from previous versions (rereferences). One paragraph.
- What is FBE, what is LDEM, what is R. Why we care about these three quantities and what is their physical meaning should be clear. Then the final products N_m and T_m , and what is their physical meaning.

$$\text{FBE}_i^{(k)} = \int dT \text{LDEM}_i(T) \text{TRF}^{(k)}(T), \quad k = 1, \dots, K \quad (1)$$

Due to unresolved coronal dynamics, tomographic reconstructions exhibit negative values of the reconstructed FBE, or zero when the solution is constrained to positive values (Frazin, 2000; Frazin, Vásquez, and Kamalabadi, 2009). These non-reconstructed voxels are indicated as black voxels in the Carington maps of the reconstructed DENT results in Section 3.

Once the LDEM is determined at each voxel, the average squared electron density $\langle N_e^2 \rangle$ and the electron mean temperature T_m in the voxel can be computed by taking its zeroth and first moments over temperature. More specifically, at the i -th voxel,

$$N_{m,i}^2 = \langle N_e^2 \rangle_i = \int dT \text{LDEM}_i(T), \quad (2)$$

$$T_{m,i} = \langle T_e \rangle_i = \frac{1}{\langle N_e^2 \rangle_i} \int dT T \text{LDEM}_i(T), \quad (3)$$

We define next a measure of the average relative difference between the tomographic and the synthetic FBE in each voxel, as

$$R_i \equiv (1/K) \sum_{k=1}^K \left| 1 - \text{FBE}_{i,\text{syn}}^{(k)} / \text{FBE}_{i,\text{tom}}^{(k)} \right|, \quad (4)$$

so that the final product of DENT is in the form of 3D maps of the electron density and mean temperature. These relationships were derived in detail in Frazin, Vásquez, and Kamalabadi 2009 (see Appendix C).

Aclarar que consideramos N_e como $\sqrt{\langle N_e^2 \rangle}$. ALBERT: Yo creo que es mejor hablar de $\sqrt{N_m^2}$

DENT has been improved since the previously published work. We now perform hollow-tomography with the EUV images, (definir hollow-tom) and a 3D regularized of the solution was applied to obtain smother and more reliable result at the edge of the tomographic grid. NOTE: NOTA de Albert: Explicá

que hacés sin justificarte. Luego aclará como se hizo eso en trabajos previos, y cual es el impacto esperado. Puedo proveer un párrafo de ejemplo, pero creo que será mejor que vos escribas como te salga y yo luego edito. Algo muy importante para todo el paper: proponete escribir “párrafos modulares”. Es decir que, con las razonables limitaciones, puedan moverse de lugar, de orden, y sigan funcionando. No definiría hollow-tomography, es un término coloquial interno nuestro. Lo correcto es que al hablar de datos mucho más arribas expliques que los datos del disco no se utilizan, pero no le des un nombre a la tomografía.

2.2. AWSoM simulations

NOTE: This section will summarize the AWSoM model. After a first version is written we will surely require some input from the Michigan friends. Points to cover:

- Data input. Link with the data input for DMT in terms of dates, etc. One paragraph.
- Main physical and technical aspects (grid in particular) of the model. One paragraph.
- Explain the artificially extended TR. One paragraph.
- What are the products of the model we care about in this paper. One paragraph.

In the transition region (TR), the plasma heats up and becomes less dense by several orders of magnitude over a short distance. Modeling this region realistically has a very high computational cost. For this reason, the model uses an artificially extended TR about $\approx 0.05 R_{\odot}$ thick, as well as artificially large values for the electronic density and temperature as boundary conditions in $r = 1 R_{\odot}$. Consequently, model results below $\approx 1.05 R_{\odot}$ were not taken into account in this article.

2.3. Tracing of results along fieldlines

NOTE: This section will explain how the AWSoM magnetic model is used to trace the thermodynamic products of both DMT and AWSoM along the fieldlines of AWSoM. This requires covering the following points:

- A first paragraph qualitatively describing the process. This includes mentioning that the AWSoM products are first interpolated into the DMT grid.
- A short paragraph describing the numerical method used to trace field lines. We have a nice version of this in your previous paper. At this point the reader should understand that for each field line we have its geometry $\mathbf{r}(s)$ as a function of the distance s along the loop, with $s(1 R_{\odot}) \equiv 0$.
- Explain very simply how then that for each voxel the middle point of the field line is selected and assigned the DMT and AWSoM products of that cell. At this point the reader should understand that for each field line we have: $Q[r(s)]$, where Q is any given quantity from either the DMT model (specifically $\sqrt{N_m^2}$ and T_m) or the AWSoM model (specifically N_e and T_e).

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- Explain fits to the DEXT results $\sqrt{N_m^2}(r)$, $T_m(r)$.
 - Explain then the classification of field lines.

For both rotations, the DEXT and AWSOM 3D maps of electron density N_e , electron mean temperature T_m and electron temperature T_e respectively were traced along the magnetic field lines modeled using AWSOM as described at the end of Section 2.1. In the case of closed loops, results were separated into their two legs, defined as the two segments that go from the coronal base up to the apex.

For each open field line, and for each leg of the closed ones, an exponential fit was then applied to the $\sqrt{N_m^2}(r)$ data points, and a linear fit applied to the $T_m(r)$ data points, as described by

$$\sqrt{N_m^2} = N_0 \exp[-(h/\lambda_N) / (r/R_\odot)] \quad (5)$$

$$T_m = T_0 + m h \quad (6)$$

where $h \equiv r - 1 R_\odot$ is the coronal height measured from the photosphere. In the electron density fit, $\lambda_N [R_\odot]$ is the density scale height and $N_0 [\text{cm}^{-3}]$ is the electron density at the footpoint ($h = 0$) of the loop. In the electron temperature fit, $m [\text{MK}/R_\odot]$ is the slope and $T_0 [\text{MK}]$ is the electron temperature at the footpoint of the loop. The slope m is the radial gradient of the electron temperature along the loop, which we denote as $m = \nabla_r T_m$ hereafter, being $\nabla_r \equiv \mathbf{e}_r \cdot \nabla$ the radial derivative operator, where \mathbf{e}_r is the heliocentric radial unit vector.

In the case of the electron density, the fitted function corresponds to the isothermal hydrostatic equilibrium solution, allowing for variation of the solar gravitational acceleration with height. This choice of function provides a straightforward means to directly evaluate how compatible is the observed coronal thermodynamical state with the hydrostatic solution.

In the present work we focus on lines with increasing and decreasing temperatures in linear form. In other words, we selected those legs where the Pearson correlation coefficient ρ between temperature T and radius r meets $|\rho(T, r)| > 0.5$. (with p-value less than 0.05). In this way, those legs in which the DEXT results cannot ensure a noticeable change in temperature between the base and the apex, are discarded.

The linear fit allows characterization of its variation with height by means of a characteristic temperature gradient $\nabla_r T_m [\text{MK}/R_\odot]$ along each leg. **En realidad esta regresion lineal utiliza de por si los errores, pero luego tmb los uso para el test de hipotesis, quizas vale la pena ponerlo solo ahi.**

To measure how good the quality of the exponential fit to the electron density and the linear fit to the temperature data are, a chi-square merit function was used (see Section 15.2 Press *et al.* (2002)). The errors were estimated in Lloveras *et al.* (2017), 0.07 MK for T_m and $4.0 \cdot 10^6 \text{cm}^{-3}$ for N_e and legs with p-value > 0.005 were rejected.

To be selected a leg must meet all following conditions.

- The leg must go through at least five tomographic grid cells with reconstructed data, and there must be at least one data point in each third of the range of heights spanned by the leg.

- ii) The value of the Pearson correlation coefficient meets $|\rho(T, r)| > 0.5$.
- iii) The p-value of both the exponential and linear fit must be less than 0.05.

For the two targeted rotations, DENT and AWSOM results were traced along the AWSOM field lines, which provides a way to characterize the global thermodynamic state of the inner solar corona and to compare them. To that end four different types of legs corresponding with four different types of magnetic structures were selected:

- Type 0 - Small (meaning with apex within the DENT range of heights) closed field lines in the latitudes range $[-30^\circ, +30^\circ]$ and $\nabla_r T_m < 0$.
- Type I - Small closed field lines in all latitudes, with $\nabla_r T_m > 0$.
- Type II - Large closed field lines (meaning their apex surpasses the DENT range of heights) in the latitudes range $[-90^\circ, -30^\circ] \cup [+30^\circ, +90^\circ]$ and $\nabla_r T_m > 0$.
- Type III - Open field lines (meaning with apex out the DENT range of heights) in the latitudes range $[-90^\circ, -60^\circ] \cup [+60^\circ, +90^\circ]$ and $\nabla_r T_m > 0$.

Hice las definiciones con gradiente de temperatura para que quede mas claro, pero en realidad uso el pearson. NOTE: Charlemos los tres en persona de esto.

2.4. Energy input flux

NOTE: This section will summarize how loop integrated energy quantities are computed from the DENT products and the AWSOM magnetic model alone. The summary will not include any derivation and will refer to Cecilia's paper. Only the results that are relevant are to be explicated. Start by explaining the model and assumptions, and end up summarizing the loop-integrated products (radiative and conductive fluxes). At the end the reader should understand that their sum is the energy input flux required at the coronal base to maintain the loop stable. This should be a very short section, as it is a summary of one previous paper alone.

Due the difference between temperatures of transition region and solar corona, there must to be a heating mechanism that compensate all the losses present in the atmosphere and maintain stable the plasma temperature. Previous works have studied this mechanism by observations and modeling, but mostly for active regions. Mac Cormack *et al.* (2017) have developed a technique that estimate the heating injected into the loops during a period of minima activity. Using DENT combined with an extrapolation of magnetic field, as we detailed before, they obtain the different plasma parameters along each magnetic loop, and therefore, compute the energy input flux required to maintain them stable. The model assume a simple energy balance for each magnetic flux tube, where the losses by radiative power (E_r) and thermal conduction power (E_c) are compensated by some coronal heating power(E_h) (Aschwanden, 2004):

$$E_h(s) = E_r(s) + E_c(s), \quad (7)$$

where s is the position along each magnetic flux tube and the powers are in units of $[\text{erg sec}^{-1} \text{ cm}^{-3}]$.

NOTE: Trabajé todo lo naranja. Primero decidí que es mejor no numerar E_c sino F_c . Eso te quedará claro cuando agregues las ecuaciones (11) y (12) que te indico más abajo y leas el último párrafo naranja de abajo, ahí cierra todo. Además describí que son los ϕ en forma correcta. El texto previo estaba muy flojón, entre otras cosas se hablaba de una (inexistente) “conductive loss function” y de una (incorrectamente denominada) “area along the loop”. Obviamente luego de chequear este texto naranja borra esta nota azul en particular.

The thermal conduction power E_c equals the divergence of the conductive heat flux F_c , i.e. $E_c(s) = [1/A(s)] d[A(s) F_c(s)]/ds$, where $A(s)$ is the transversal area of the magnetic flux tube at position s . Under a quiescent solar corona plasma regime, the conductive flux is assumed to be dominated by the electron thermal conduction, described by the usual Spitzer model (Spitzer, 1962)

$$F_c(s) = -\kappa_0 T(s)^{5/2} \frac{dT}{ds}(s), \quad (8)$$

where $\kappa_0 = 9.2 \times 10^{-7} \text{ erg sec}^{-1} \text{ K}^{-7/2}$ is the Spitzer thermal conductivity.

Radiative power depends on the amount of plasma in certain temperatures that can radiate. The model estimate it by integrating the squared density multiplied by a radiative loss function $\Lambda(T)$. This function depends on plasma temperature and is calculated by the atomic database and the plasma emission model from CHIANTI (Del Zanna *et al.*, 2015). Then, the expression for radiative power obtained is:

$$E_r = \int dT L_{\text{DEM}}(T) \Lambda(T) \quad (9)$$

The energy balance given by Equation (7) is then integrated in the volume of any given coronal magnetic flux tube. Dividing the result by the flux tube area at the coronal base, and making use of the soleidonal condition of the magnetic field, a field-line integrated version of that energy balance is found,

$$\phi_h = \phi_r + \phi_c. \quad (10)$$

where the line integrated flux quantities $\phi_{r,c} [\text{erg sec}^{-1} \text{ cm}^{-2}]$ are given by (Mac Cormack *et al.*, 2017),

NOTE: Diego, please include here Equations (11) and (12) of that paper, which incredibly enough will have the same numbers here.

Note that, for any given field line, all quantities in these two expressions can be computed from the DEMT and AWSOM models through Equations (8)-(9). Once computed, the quantity ϕ_h can be calculated, which is the energy input flux required at the coronal base of each coronal field-line required to maintain a stable coronal structure.

3. Results

3.1. Tomographic Results

NOTE: This section will show first the DENT products in the form of Carrington maps at three heights, for the two target rotations. It will then show the latitude/longitude location of the four types of magnetic loops (defined in Section 2.3) at a sample height. Then, for the each type of loop, we will show histograms of the values of different DENT products traced along them, characterizing in a quantitative fashion the DENT results in distinct magnetic structures. It will also include a quantitative comparison between the two target rotations of the median values of the DENT products in each type of loop.

To perform a comparative study of the thermodynamic structure of the inner corona in different solar cycles we carry out a DENT reconstruction for two rotations: CR-2082, using STEREO/EUVI data and CR-2208, using SDO/AIA data. Once the LDEM was determined for each rotation, the square root of the mean value of the electron density squared $\sqrt{N_m^2}$ and the electron mean temperature T_m were computed at each voxel of the tomographic computational grid by means of Equations 2 and 3. As an example, Figures 1 and 2 show the latitude-longitude maps of DENT results for both rotations at three different heights of interest. The lowest height where the tomography gives results, the lowest height where the AWSOM model results are reliable and a representative average height were selected, that is $r = 1.025$, 1.065 and 1.105 , R_\odot respectively. Black voxels correspond to non-reconstructed voxels (explained in Section 2.1). Thick-black curve indicates the magnetic open/closed boundary based on the magnetic field of AWSOM model.

While CR-2082 was highly axisymmetric and showed two ARs (active region) in latitude $+30^\circ$ and longitude 50° and 120° . CR-2208 showed two AR, one around latitude $+5^\circ$ and longitude 140° and the other around latitude-longitude $+5^\circ$, 300° . Except for these ARs, in both rotations the open-closed boundaries derived from the respective AWSOM model match the shape of the contour levels of the tomographic maps of both the electron density and the electron mean temperature. In the AR, there is a lack of match between the open-closed boundaries and the tomographic reconstructions of density and temperature, that is why those ranges were excluded from the analysis.

PARRAFO CORTO DE OBSERVACIONES QUE LUEGO SERAN CORROBORADAS CON HISTOGRAMAS. A visual inspection of Figures 2 and 3 suggest/indicates that the streamer region of CR-2082 was denser and colder than CR-2208. In both rotations, but specifically in CR-2082 the temperature in the equatorial latitudes tend to decrease with height while the latitudes between the open-close boundaries and the equatorial latitudes tend to increase. These thermodynamic structures have been reported for other rotations in previous works (Lloveras *et al.*, 2017). Closed regions are characterized by relatively larger values of density and temperature in comparison to the open regions. **que mas se puede agregar?**

To study the variation of the DENT results in different magnetic structures we traced $\sqrt{N_m^2}$ and T_m along the magnetic field lines of the AWSOM model.

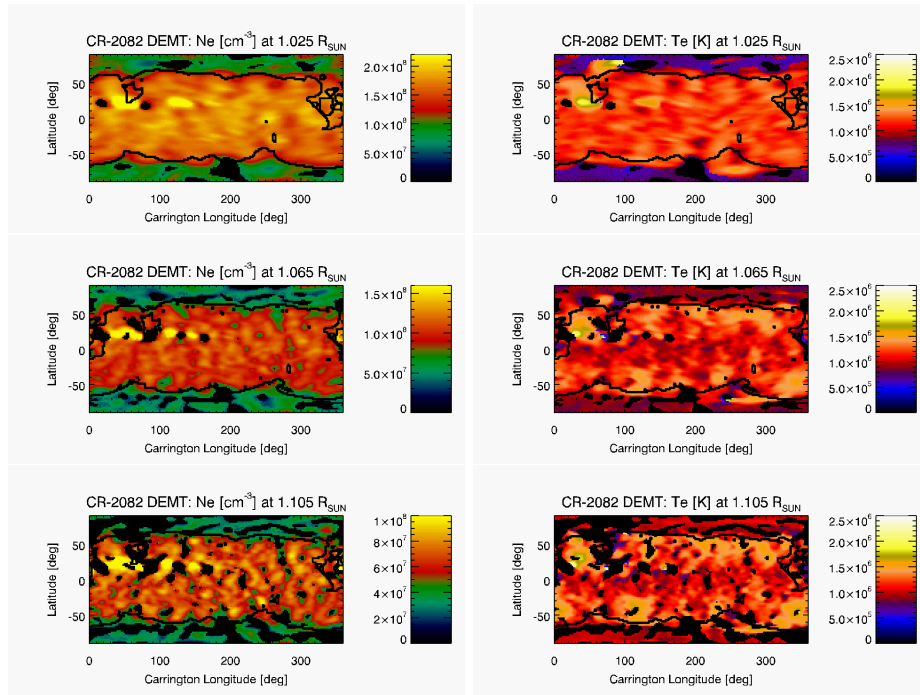


Figure 1. Carrington maps of DMT products $\sqrt{N_m^2}$ (left panels) and T_m (right panels) for CR-2082. Top, middle and bottom panels show the results at three heliocentric heights, 1.025, 1.065 and 1.105 R_\odot respectively. Black voxels correspond to non-reconstructed regions (see text in Section 3.1) and thick-black curves indicate the open/closed boundaries.

For both rotations, all field lines that meet the criteria listed in Section 2.3 were selected. For each field line, the data points of electron density and electron mean temperature as a function of height were fitted to the Equations 5 and 6.

As a result, the electron density $N_0(r = 1.0 R_\odot)$ and scale height λ_N were computed for each leg, as well as the temperature gradient $\nabla_r T_m$, and the height-averaged (along the leg) electron temperature $\langle T_m \rangle$. As the DMT technique studies only coronal conditions, *i.e.* above height $r \approx 1.025 R_\odot$, the exponential fit to the electron density data as a function of height may be used to compute the tomographic electron density at the coronal base $\sqrt{N_m^2}(r = 1.025 R_\odot)$ as in previous articles. In this case, for further comparison analysis with the AWSOM results, we selected the lowest height where the results of the MHD model are reliable, that is 1,055 R_\odot .

As a way of knowing the physical position of all traced lines, Figure 3 shows the latitudinal-longitudinal location of all open and closed file lines traced in gray scale at 1.105 R_\odot for both rotations, where the ARs were excluded. At top, all traced legs with enough tomographic data are shown. In black all legs corresponding to closed lines and in grey corresponding to open lines. This traced lines were divided into four types as explained in Section 2.3 and the fit in density and temperature was applied, selecting only those legs that met the criteria.

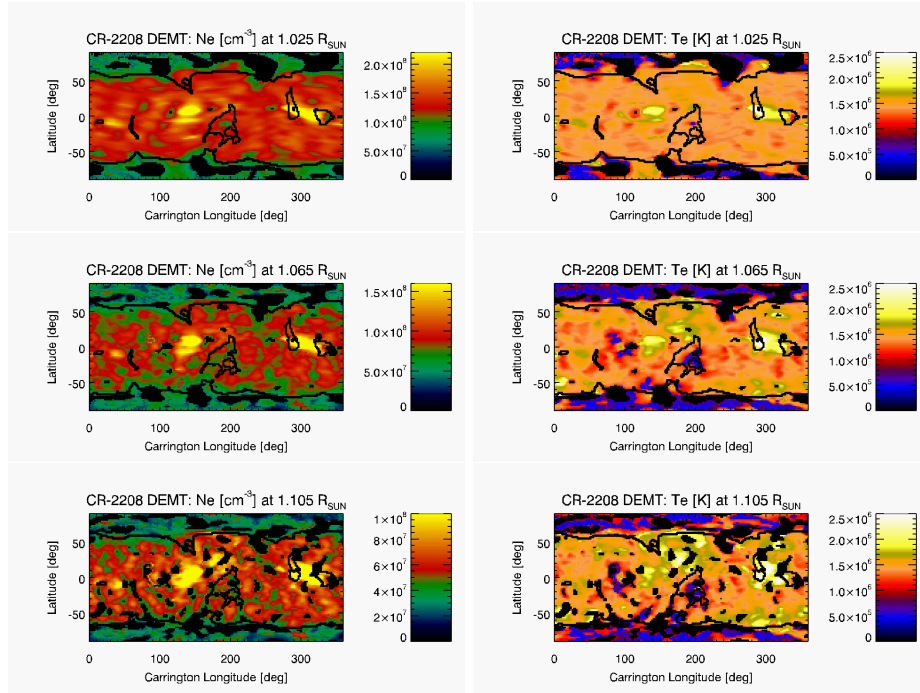


Figure 2. Same as Figure 1 but for CR-2208.

The bottom figure shows the latitudinal-longitudinal location of this selected lines at same height. In blue the type 0 lines are indicated, populating the most equatorial latitudes (dominated mostly by down legs). This kind of structures has been informed in previous works, Lloveras *et al.* (2017) observed down legs in a rotation during the solar minima of 1996 and for the previous rotation CR-2081. A decrease in the amount of down legs in equatorial latitudes could be an indicator of increased coronal activity Nuevo *et al.* (2013). In red the type I lines, populating all the streamer but mostly the mid latitudes. In violet the type II lines, which are mostly very large trans-equatorial field lines forming the envelope of the Streamer belt, and in green dots the type III lines, populating the coronal holes.

Of the 45000 legs selected in 2082, 14% are type 0, 38% are type I, 23% are type II and 25% type III, on the other hand of the 59000 legs selected in 2208, 8% are type 0, 43% are type I, 22% are type II and 30% type III.

This selection of types of lines automatically leads to specific regions arranged along all longitudes associated with different thermodynamic, with good symmetry between both hemispheres in both rotations.

To compare obtained traced results in both rotations and for each type of lines (associated with different types of regions), Figure 4 shows histograms (normalized to the total amount of the field lines selected) of the temperature radial gradients ($\nabla_r T_m$). The lack of population around values close to zero is due to the requirement $|\rho_t| > 0.5$ which discards quasi-isothermal legs.

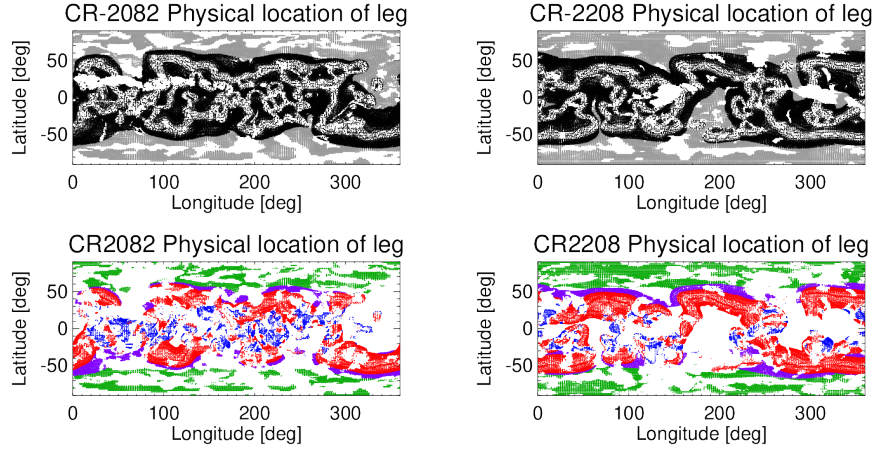


Figure 3. *Top panels:* latitude-longitude location at heliocentric height $r = 1.105 R_{\odot}$ of all open (grey) and closed (black) traced field lines for which there are along at least five DENT data points, for both CR2082 (left) and CR2208 (right). *Bottom panels,* latitude-longitude location (at the same height) of each traced leg that meets the criteria listed in Section 2.3. Type 0, I, II and III legs (see text) are shown in blue, red, violet and green, respectively, for CR-2082 (left) and CR-2208 (right).

Both rotations are characterized by lines of type 0 with mean values of $-2.5 \text{ MK}/R_{\odot}$, mean values of $2.3 \text{ MK}/R_{\odot}$ in type I and $2.4 \text{ MK}/R_{\odot}$ in type II. The only notable difference is the smaller number of type 0 lines in CR-2208 and the mean value of $4.5 \text{ MK}/R_{\odot}$ in type III for CR-2208 which is twice that for the other rotation.

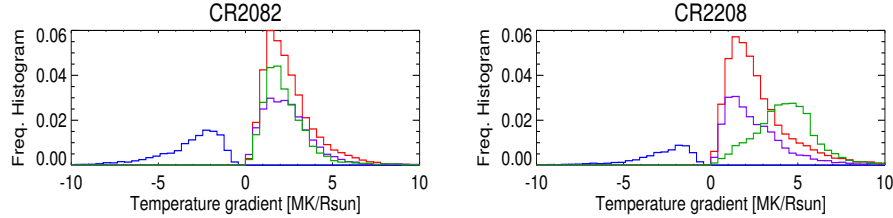


Figure 4. Frequency histograms of the temperature radial gradient for the different four types of legs shown in Figure 3 (using the same color code) for CR-2082 (left panel) and CR-2208 (right panel).

As an example of the traced magnitudes, Figure 5 shows the statistical results of both rotations for all field lines types. From top to bottom type 0 to type III. From left to right the panels show the statistical distribution of $N_{CB} \equiv \sqrt{N_m^2}(r = 1.055)$, λ_N , $\langle T_m \rangle$, in each panel the median m is indicated.

For all type of lines, Table 1 summarize the quantitative comparative analysis between the results of the two target rotations sampling both minima. For CR-2082 quantities are expressed as absolute values, while for CR-2208 they

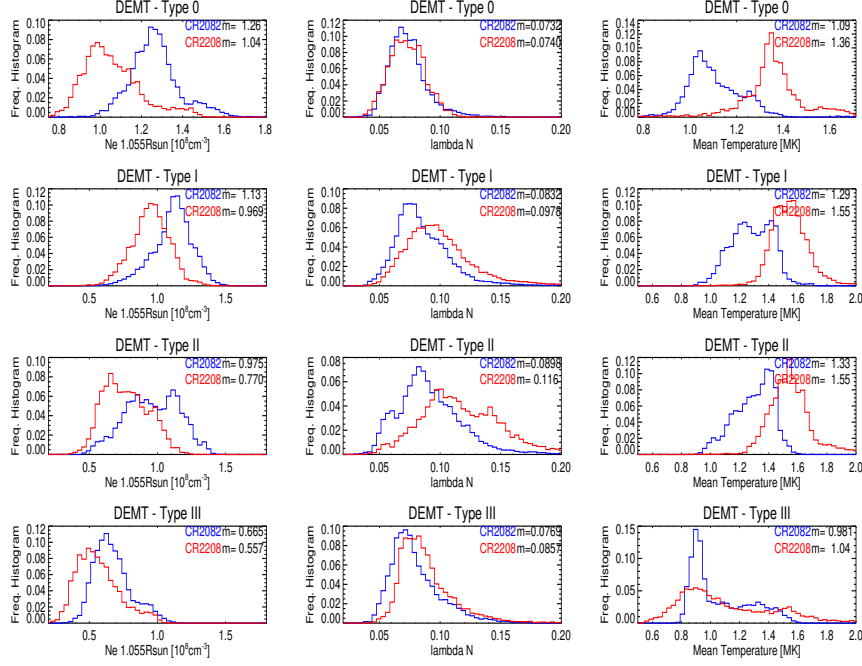


Figure 5. Statistical distribution of DENT results traced along legs of type 0, I, II and III (from top to bottom), as defined in Section 2.3. From left to right: electron density at the coronal base $N_{CB} \equiv \sqrt{N_m^2}(r = 1.055 R_\odot)$, electron density scale height λ_N , and loop-averaged electron temperature $\langle T_m \rangle$. Results for target rotations CR2082 and CR2208 are shown in blue and red color, respectively. In each panel the median value m is indicated. **NOTE:** (by Albert) The histogram figures in this draft need to be improved, in particular font size-and-type need to be optimized for clear reading. Also, we need to improve their X/Y ratio, here I stretched the Y size a bit in the LaTeX source. Also, I believe you should use a FIXED x-range for each quantity (one for N_{CB} , one for λ_N and one for $\langle T_m \rangle$) across all Figures 5, 12, and 13, as that will visually highlight how medians and standard deviations compare, both between targets (CR2082 versus CR2208, this Figure) as well as between models (DENT versus AWSOM, Figures 12 and 13).

are informed relative to the corresponding results for CR-2082. The following major results concerning the structure of each rotation individually and their comparison can be drawn.

si queda una asociacion directa entre los tipos de lineas y las estructuras de streamer, CH etc entonces usaria esos nombres a continuacion.

From the equatorial streamer of both rotations, increasing latitudes show decreasing coronal base density, and increasing density scale height and electron temperature. In both rotations the CH regions are characterized sub-MK temperatures, and coronal base densities of order $\approx 1/2$ of the equatorial streamer. Histograms with sharp populations indicates that for both rotations the values of all physical quantities exhibit a strong symmetry between both hemispheres.

A comparison of the results between the two rotations shows that CR-2208 is characterized by lower values (between 11% and 16%) of coronal base electron

Table 1. Median value (indicated as “Md”) of the statistical distribution of N_{CB} , λ_N , and $\langle T_m \rangle$ for each coronal type of lines defined in Section 2.3. For CR-2082 values are expressed in absolute terms, while for CR-2208 they are informed as a percentual variation relative to the CR-2082 value.

Type	Md(N_{CB}) [10^8 cm^{-3}]	Md(λ_N) [10^{-2} R_\odot]	Md($\langle T_m \rangle$) [MK]
0	1.26 (−17.4%)	7.3 (+13.6%)	1.09 (+24.7%)
I	1.13 (−14.1%)	8.3 (+18.1%)	1.29 (+20.1%)
II	0.97 (−20.6%)	8.9 (+29.1%)	1.33 (+16.5%)
III	0.66 (−16.2%)	7.7 (+11.4%)	0.98 (+ 6.1%)

density than CR-2082, higher values of density scale height (between 4% and 19%) and higher values of electron temperature (between 6% and 21%).

To compare integrated quantities along loops, we select those close loops where both legs meet the requirements of Section 2.3. In an analogous way to the legs definition, we will refer to them as loops of type 0, I or II to those whose both legs meeting the criteria. When two legs of the same loop have a decreasing temperature with height (type 0), they are called down loops.

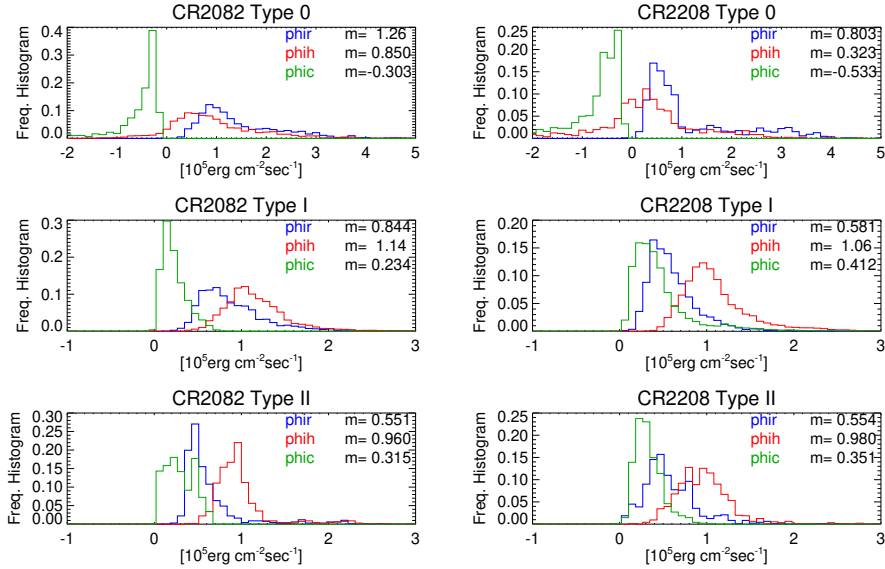


Figure 6. Statistical results of the loop-integrated energy flux quantities ϕ_r , ϕ_c , and ϕ_h in colors blue, red and green, respectively for CR-2082 (left) and CR-2208 (right). From top to bottom, panels show the results for loops of type 0, I and II, which are loops for which both legs meet the criteria from Section 2.3. NOTE: (by Albert) En el XYOUTS podés poner la letra griega ϕ usando los comandos “!4” y “!3”, y para que te salga r,c,h como subíndice pero no tan pequeño usá: “!Dp!N”, dónde p=r,c,h.

For both target rotations, and for loops of type 0, I and II, Figure 6 shows the frequency histogram of the loop-integrated energy flux quantities ϕ_h , ϕ_r and ϕ_c in green, blue and red color, respectively.

Type 0 loops present slightly larger gradients than type I. Therefore, type 0 present bigger amount, in modulus, of loop-integrated conductive flux. Due conductive flux negative can not be assumed as a loss but it can be contributing to the coronal heating. Thus, to compensate the radiative losses, a lower loop-integrated energy input flux (ϕ_h) is required.

We can see, on type 0 loops, a median value of energy input flux for CR2082 twice bigger than CR2208. Also, in both rotations, we observe a negative population of loop integrated energy input flux. This can refer to some loops where loop-integrated radiative flux is too small in comparison with the larger gradients of negative conductive flux, resulting in negative energy input flux.

The type I loops show a loop-integrated positive conductive for each loop in both rotations consistent with positive gradients of temperatures. The conductive flux have values around $0.3 \cdot 10^5 [\text{erg sec}^{-1} \text{cm}^{-2}]$ in both rotations. Similar results have been found for loop-integrated radiative flux and energy input flux for both rotations.

No son los mismos colores que el paper de ceci, tmb quizas podriamos poner la cant de LOOPS y los valores medianos, ceci insiste.

3.2. Tomography and Models

NOTE: This section will show the results of AWSoM for N_e and T_e , and how they compare with the corresponding results of DMT ($\sqrt{N_m^2}$ and T_m). It will first show the Carrington maps of N_e and T_e at the same three heights we showed the same maps for the DMT results. Then it will show the lat/lon location of loops of type I, II and III at a middle height, just loke for DMT results. As the two targets are highly axis-symmetric, a first quantitative comparison between AWSoM and DMT will be shown as longitude-averaged latitudinal profiles of the results of both models at a sample height. To make comparisons in distinct magnetic structures, we will also show the average radial dependence of results for loops of type I, II and III, overplotting the AWSoM and the DMT results. For the tree types of loops we will show histograms of the DMT and AWSoM products along field lines. This is a LOT of information from which and we need to pull out a few clear conclussions.

The thermodynamic results of AWSoM are shown in Figures 7 and 8. That is latitude-longitude maps for N_e and T_e modeled for CR-2082 and CR-2208 at $r = 1.025, 1.065$ and $1.105 R_\odot$. The black curve denotes the same open-close boundary showed in Figures 1 and 2 respectively.

A visual inspection of these maps shows that both reconstructions are very axysymmetric for N_e and T_e as well. The closed region is denser and hotter than the open region. The density and temperature gradients are maximum in the open-close boundary. The open field regions in low latitudes maintain thermodynamic values associated with open regions. The reconstruction of CR-2082 does not show clear the ARs, while CR-2208 shows a well reconstruction of both AR around latitude-longitud $+5^\circ, 140^\circ$ and $+5^\circ, 300^\circ$. In both rotations

the open-closed boundaries derived from the AWSoM model match the shape of the contour levels of both the electron density and the electron temperature.

The same intensity scales in both density and temperature can be observed between the tomographic results and those of AWSoM in both rotations, as well as the same thermodynamic structures. That is, equatorial latitudes, mid latitudes and the open regions.

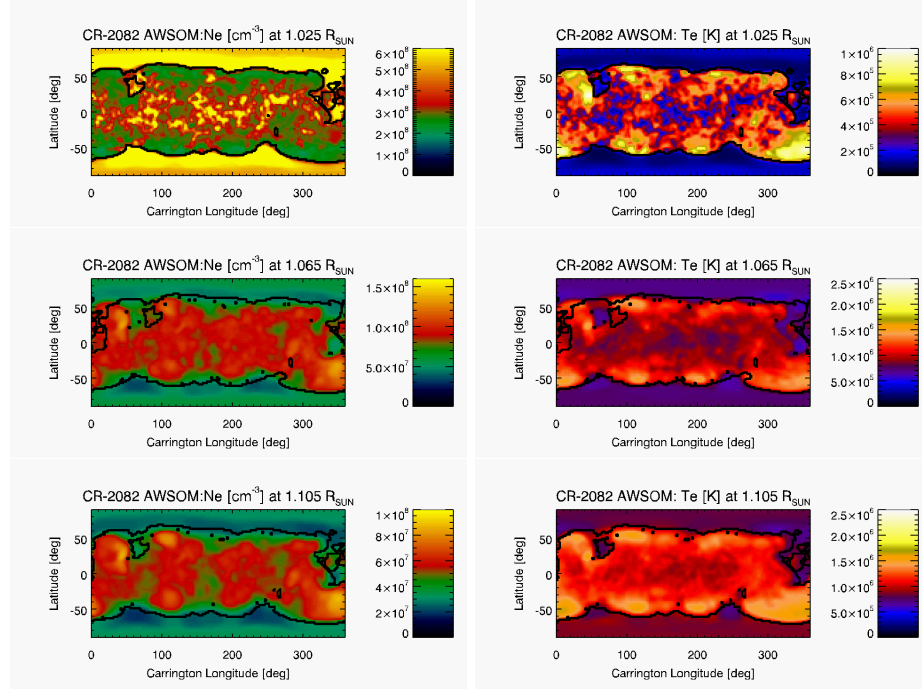


Figure 7. Carrington maps of density (left panels) and temperature (right panels) obtained with AWSoM model for the same three heights shown in Figure 1.

In the same way as in Section 3.1, the model results were traced along the magnetic lines obtained with it. This allowed to obtain $N_e(r)$ and $T_e(r)$ along the magnetic lines and fit them to the Equations 5 and 6.

The actual version of the AWSoM model can not reproduce down legs, so for a further comparison we only select lines of type I, II and III of those meeting the criteria described in Section 2.3.

As a way to understand the distribution of legs, Figure 9 shows the latitudinal-longitudinal location of the traced legs at top and the selected fitted legs meeting the criteria at $1.105 R_{\odot}$ for both rotations, in the same way as in Section 3.1. In red the type I lines, populating mostly the entire streamer. In violet the type II lines, populating the edge that assemble the envelope around the streamer, and in green dots the type III lines, populating the coronal holes.

The reader may notice that these structures are the same as those in Figure 3 but more filled. The cause is twofold, on the one hand the same magnetic model

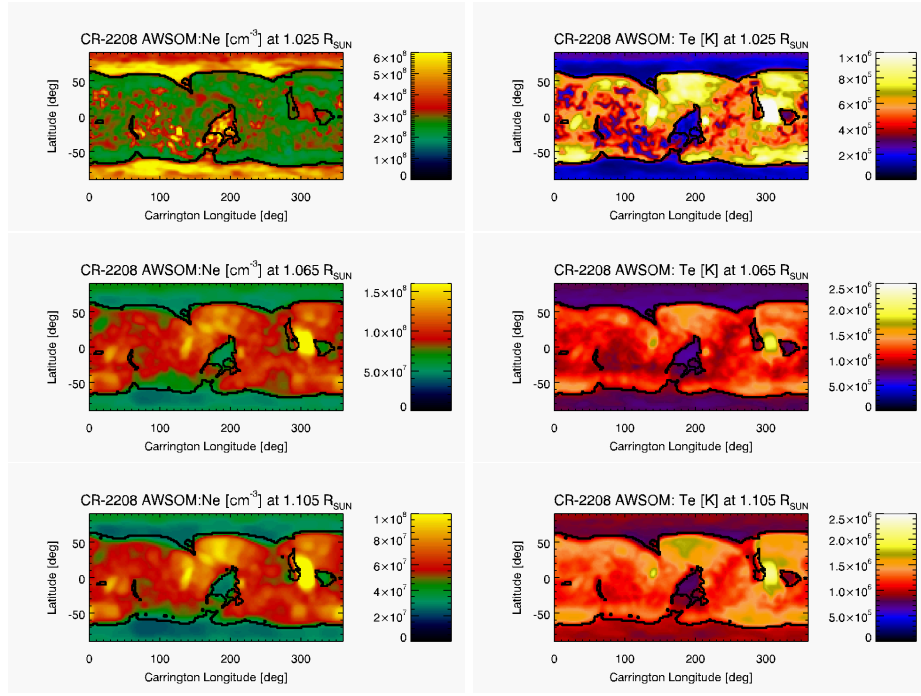


Figure 8. Same as Figure 8 for CR-2208.

was used to trace both DENT and AWSOM results, on the other hand is because giving the model smoother results, fewer fitted lines are filtered.

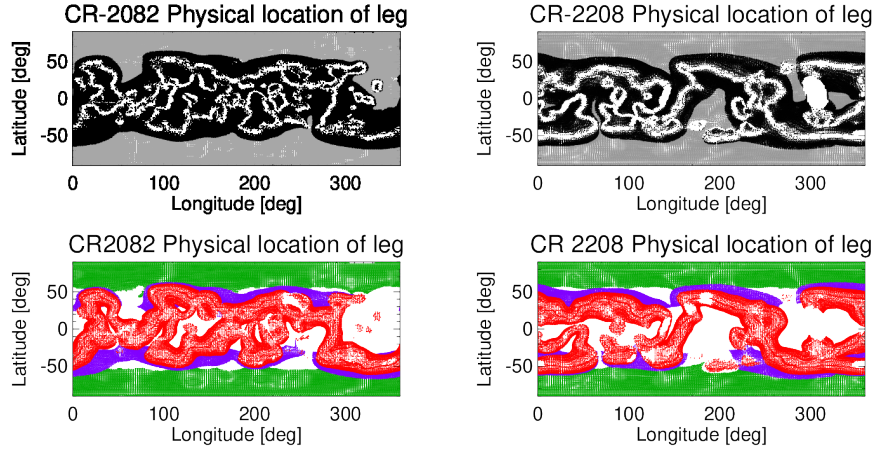


Figure 9. Same as Figure 3, but using the density and temperature of the AWSOM model to classify its legs in types I, II and III. The model does not exhibit legs of type 0.

To highlight the latitudinal structures observed in Figures 1 and 7 for CR-2082 as well as Figures 2 and 8 for CR-2208, Figure 10 shows for both rotations the latitudinal variation of both the electron density and the electron mean temperature at the height $r = 1.105 R_{\odot}$, averaged over all longitudes (excluding the AR). DENT results are shown in blue lines while AWSOM results in red. While the temperature shows a maximum in each hemisphere, the density has a simpler behavior, with higher values in the low-latitudes of the streamer and a nearly monotonic decrease towards larger latitudes. The vertical black lines denotes the average of the open-close boundary in both hemispheres to each rotation.

Me gustaria hacer un comentario de cierre para esta figura que refiera las Figuras 3 y 9, como la linea de abajo, y que conecte con los histogramas que vienen a continuacion. Si no queda bien, hay que mover esta figura entre la actual figura 8 y 9.

This longitude-averaged behaviors match well with the physical locations of the legs shown in Figures 9 and 3.

According to the visual inspection on the maps of carrington, the latitudinal profiles comparing both DENT and AWSOM results are on the same scale. Little difference in density is observed in CR-2208 and in temperature in CR-2082, while DENT in CR-2082 was 20% denser. Finally, DENT results for CR-2208 were 25% hotter in the streamer and 30% colder in the polar zone than AWSOM. Density profiles on AWSOM show an increase around the polar zone, which seems like an artifice.

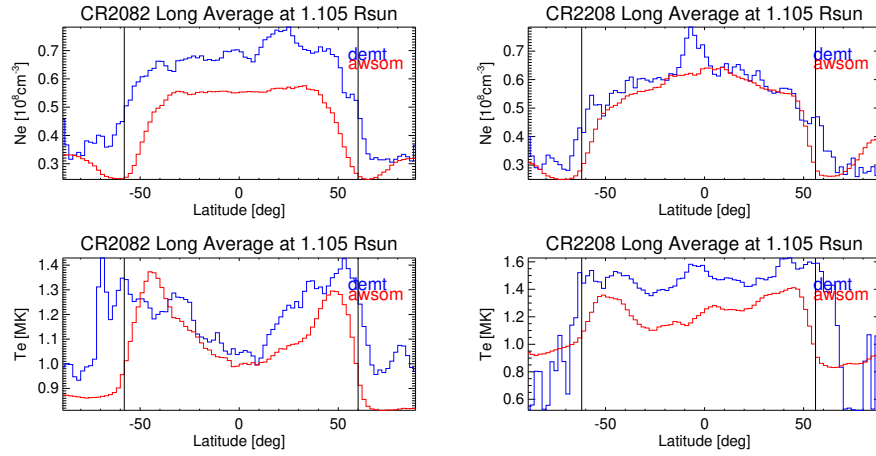


Figure 10. Longitude-averaged latitudinal dependence of the electron density (top) and temperature (bottom) for DENT (blue) and AWSOM (red) results at $1.105 R_{\odot}$. The left (right) panels correspond to CR-2082 (CR2208). The vertical black line indicates the longitude-averaged latitudes of the open/close magnetic boundary in both hemispheres. **NOTE: (by Albert):** The idea of this Figure is to show how DENT and AWSOM compare AS A FUNCTION OF LATITUDE at an intermediate height of the DMET grid, averaged in longitude (excluding zones with ARs).

To summarize the comparison between both results during both rotations at full height, it is shown in Figure 11 the average fits of $N_e(r)$ and $T_e(r)$ for legs of type I (blue), II (red), and III (green) for both rotations.

In average all temperature and density behavior from both results are on the same scale. The coronal base densities of AWSoM were lower than those of DENT in CR-2082 and higher in CR-2208. The fact that the density decays to similar values between both results suggests a similar height scale. On the other hand basal temperature of AWSoM was lower in both rotations, but the temperature gradients are very similar in each type of legs, showing a very similar increase of temperature with height.

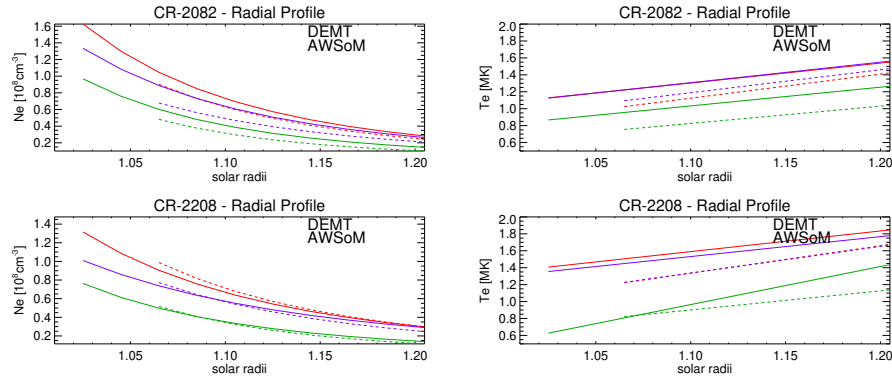


Figure 11. Average fits to $N_e(r)$ (left panels) and $T_e(r)$ (right panels) for legs of type I (red), II (violet), and III (green), for CR-2082 (top panels) and CR-2208 (bottom panels). Solid lines correspond to DENT results while dashed lines correspond to AWSoM results. **NOTE: (by Albert):** The idea of this Figure is to show how DENT and AWSoM compare AS A FUNCTION OF HEIGHT, in three distinct types of magnetic structures, and for both rotations.

To summarize the statistical comparison, Figure 12 and 13 shows the statistical results of both rotations for the three types of field lines traced with DENT (blue) and AWSoM results (red). From left to right the panels show the statistical distribution of N_{CB} , λ_N and $\langle T_m \rangle$. For all types of lines, Table 2 summarizes the median value of the statistical distribution of each physical quantity derived from the analysis. The DENT quantities are expressed as absolute values, while for AWSoM they are informed relative to the corresponding result for DENT.

aca podria ir un pequeño parrafo que resuma las diferencias. Pero justo a continuacion empiezan las conclusiones y uno tiende a repetir, no se como encararlo para que no quede dos veces lo mismo y tan pegado uno de lo otro.

Figure 14 shows the longitude-averaged AWSoM radial wind speed V_r at $6 R_\odot$, where all field lines are open. The heliocentric current sheet (HCS) location is indicated by the minimum of the speed curve.

can we say something without having traced the lines? **NOTE: (by ALBERT)** Todo el punto de la última Figura es precisamente decir algo. La comparación con la Figura 10 muestra una anti-correlación entre la densidad basal de DENT y la velocidad terminal de AWSoM.

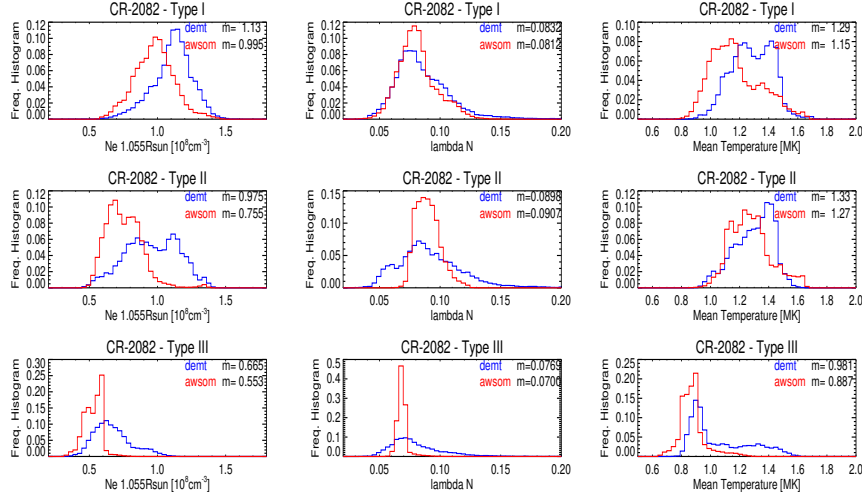


Figure 12. Statistical distribution of DMT (blue) and AWSOM (red) results traced along legs of type I, II and III (from top to bottom), as defined in Section 2.3. From left to right: electron density at the lowest coronal height of the AWSOM model $N_e(r = 1.055 R_\odot)$, electron density scale height λ_N , and leg-averaged electron temperature $\langle T_m \rangle$. In each panel the median value m is indicated.

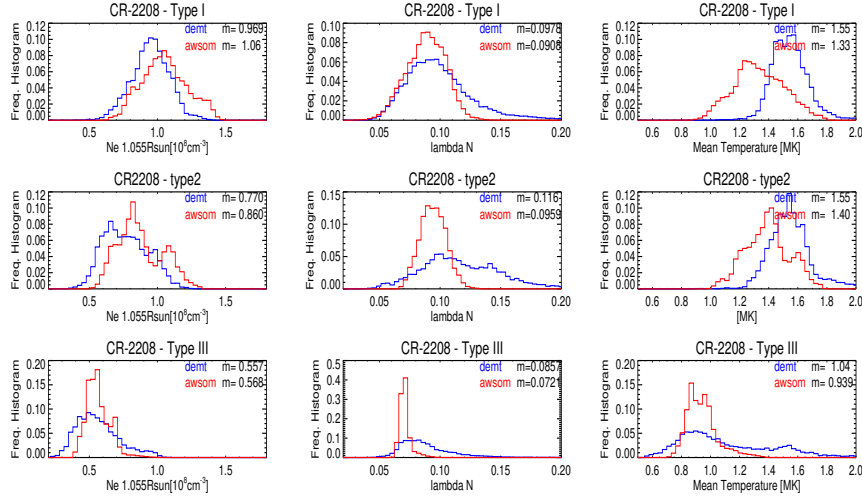


Figure 13. Same as Figure 12 for CR-2208.

For each rotation, everything to the south of the HCS maps down to the southern open region in the Figures 10, and the same for the northern hemisphere, showing an anti-correlation between N_e at low heights and high V_r at $6 R_\odot$.

Table 2. Median value (indicated as “Md”) of the statistical distribution of N_{CB} , λ_N , and $\langle T_m \rangle$ for each coronal type of lines defined in Section 2.3. DENT values are expressed in absolute terms, while AWSoM results are informed as a percentual variation relative to the corresponding DENT value.

Type	Md(N_{CB}) [10^8 cm^{-3}]	Md(λ_N) [$10^{-2} R_\odot$]	Md($\langle T_m \rangle$) [MK]
CR-2082			
I	1.13 (-11.9%)	8.3 (-2.4%)	1.29 (-10.8%)
II	0.97 (-22.5%)	8.9 (+2.2%)	1.33 (-4.5%)
III	0.66 (-16.8%)	7.7 (-9.1%)	0.98 (-9.2%)
CR-2208			
I	0.97 (+9.3%)	9.8 (-7.1%)	1.55 (-14.2%)
II	0.77 (+11.7%)	11.6 (-17.2%)	1.55 (-9.7%)
III	0.56 (-1.9%)	8.6 (-16.3%)	1.04 (-9.6%)

Voy a intentar jugar con las lineas tipo 3 aumentando su cota inferior de latitud para ver si los histogramas de temp media dejan de ser tan dispersos y se vuelven mas sub-1MK.

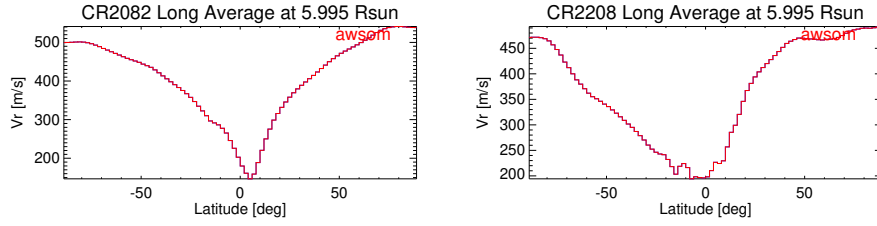


Figure 14. Longitude-averaged latitudinal dependence of the AWSoM model wind speed at $6.0 R_\odot$ for CR2082 (left panel) and CR2208 (right panel). NOTE: (by Albert): Even if this is not yet the terminal speed, it is a pretty high height, and one can already see the transition from FAST to SLOW speed in the model, where the absolute minimum indicates the location of the HCS at this height. I believe it is interesting to compare these two curves with Figure 10. Note that while in this Figure all latitudes are magnetically open, in Figure 10 only the part at latitudes larger than the vertical black lines are open. Note that DENT shows that in the open low corona N_e decreases from the O/C boundary towards the poles, anti-correlating with the AWSoM terminal speed, which is nice. In order to quantify this anti-correlation we need to trace the terminal speed and the DENT N_e , which we may need to do with the help of Chip and Nishtha. Chip, we can discuss this idea over Skype after Xmas if you are around.

4. Discussion

NOTE: (by Albert) This section will discuss and summarize some main conclusions, possibly in three groups: a) DENT analysis alone (including thermodynamical structure, up/down loops, energy considerations, etc.), b) Comparison AWSoM/DENT (longitude-averaged latitudinal profiles, height-dependence

profiles in different magnetic structures, etc), and c) Correlation between AW-SoM terminal wind speed and DEMT basal density, I will expand on this later this week as this may require help from Michigan).

NOTE: Here is a paragraph describing the empirical successes in the reconstruction of both rotations.

*In comparing the DEMT results obtained for the two selected targets, it is important to bear in mind they rely on data provided by two different instruments: AIA and EUVI. In order to quantify the systematic difference of the DEMT products based on both instruments, Nuevo *et al.* (2015), who were the first to apply DEMT to AIA data, analyzed a single target using both instruments independently. They found that while the density product is essentially equal, the temperature product of DEMT based on AIA data is systematically 8% larger than the one based on EUVI data, i.e. $T_m^{(AIA)}/T_m^{(EUVI)} \approx 1.08$. Considering such correction, Figure 5 and Table 1 indicate that CR-2208 was $\approx 10 - 15\%$ hotter than CR-2082 throughout the streamer belt region. As for the electron density products, CR-2208 was found to be $\approx 15 - 20\%$ less dense than CR-2082 throughout the streamer belt region. These systematic differences are beyond the uncertainty level in the DEMT products due to systematic sources (radiometric calibration and tomographic regularization), that Lloveras *et al.* (2017) estimated to be $\lesssim 5\%$ and $\approx 2\%$ for the electron temperature and density, respectively.

*Coronal magnetic structures for which the temperature increases/decreases with height in the range of heights covered by DEMT, $1.02 - 1.22 R_\odot$, have been dubbed as “up”/“down” loops by Huang *et al.* (2012) and Nuevo *et al.* (2013), who first observed their presence by means of DEMT. As speculated by the authors of those works, loops of type down can be expected if the heating deposition is strongly confined near the coronal base of a magnetic loop. Down loops were first predicted by Serio *et al.* (1981), and later on by Aschwanden and Schrijver (2002). In a recent study by Schiff and Cranmer (2016), down and up loops have been successfully reproduced by a numerical implementation of a 1D steady-state model that considers time-averaged heating rates.

We found down legs in the DEMT reconstructions mainly in the equatorial latitudes, being a greater quantity and better distribution along longitudes in CR-2082 than in CR-2208. Nuevo *et al.* (2013) showed a decrease in the amount of down loops as the solar activity increases, which is consistent with the results shown in this article.

*We can see, on type 0 loops, a median value of energy input flux (ϕ_h) for CR-2082 twice bigger than CR-2208. In both rotations this magnitude has negative values, this can refer to some loops where loop-integrated radiative flux ϕ_r is too small in comparison with the larger gradients of negative conductive flux ϕ_c , resulting in negative energy input flux.

This effect can be due a subestimation of the density present in the loop or the limited range of temperature used to reconstruct tomographic parameters, that can affect the compute of radiative power. Nevertheless, this population represent lower than the **ALGÚN PORCENTAJE%** of total population in both cases.

Consistent that the density in CR-2082 was higher, the ϕ_r was also higher. In both rotations the magnitudes in the temperature gradients are similar, but the temperatures in CR-2208 were consistently higher at all heights, resulting in a larger ϕ_c in module. This results are in agree with the observations of Mac Cormack *et al.* (2017).

THE AR of CR-2082 wasn't well modeled by AWSoM.

In both rotations, AWSoM and DENT shows same thermodynamic structures along latitude and longitude with very comparable values between $1.055 - 1.25 R_\odot$

Figure 12 and Table 2 indicate that AWSoM reconstruction of CR-2082 was found to be less dense $\approx 11 - 22\%$ and colder $\approx 5 - 10\%$.

Both reconstructions of CR-2208 reconstructed the ARs with very similare values and in almost the same latitude-longitude location. Figure 13 and Table 2 indicate that AWSoM reconstruction of CR-2208 was found to be denser $\approx 0 - 10\%$ and colder $\approx 10 - 15\%$.

CR-2208 was $\approx 10 - 15\%$ hotter than CR-2082 throughout the streamer belt region. As for the electron density products, CR-2208 was found to be $\approx 15 - 20\%$ less dense than CR-2082 throughout the streamer belt region. These systematic differences are beyond the uncertainty level in the DENT products due to systematic sources (radiometric calibration and tomographic regularization), that Lloveras *et al.* (2017) estimated to be $\lesssim 5\%$ and $\approx 2\%$ for the electron temperature and density, respectively.

I would like to say something about FFigure 10. Maybe the weird tale in AWSoM Ne in the poles. Ideas?

We could say something about the anti correlation between speed and density and not get into the trace because issues do always appear and we don't have much time, and we all want to be happy in january. Unless the collaborators find easy to make this part.

References

- Aschwanden, M.J.: 2004, *Physics of the Solar Corona. An Introduction*, Praxis Publishing Ltd, ??? ADS.
- Aschwanden, M.J., Schrijver, C.J.: 2002, Analytical Approximations to Hydrostatic Solutions and Scaling Laws of Coronal Loops. *Astrophys. J. Suppl.* **142**(2), 269. DOI. ADS.
- Del Zanna, G., Dere, K.P., Young, P.R., Landi, E., Mason, H.E.: 2015, CHIANTI - An atomic database for emission lines. Version 8. *Astron. Astrophys.* **582**, A56. DOI. ADS.
- Frazin, R.A.: 2000, Tomography of the Solar Corona. I. A Robust, Regularized, Positive Estimation Method. *Astrophys. J.* **530**, 1026. DOI. ADS.
- Frazin, R.A., Vázquez, A.M., Kamalabadi, F.: 2009, Quantitative, Three-dimensional Analysis of the Global Corona with Multi-spacecraft Differential Emission Measure Tomography. *Astrophys. J.* **701**, 547. DOI. ADS.
- Huang, Z., Frazin, R.A., Landi, E., Manchester, W.B., Vázquez, A.M., Gombosi, T.I.: 2012, Newly Discovered Global Temperature Structures in the Quiet Sun at Solar Minimum. *Astrophys. J.* **755**, 86. DOI. ADS.
- Lloveras, D.G., Vázquez, A.M., Nuevo, F.A., Frazin, R.A.: 2017, Comparative Study of the Three-Dimensional Thermodynamical Structure of the Inner Corona of Solar Minimum Carrington Rotations 1915 and 2081. *Solar Phys.* **292**(10), 153. DOI. <https://doi.org/10.1007/s11207-017-1179-z>.

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- Mac Cormack, C., Vázquez, A.M., López Fuentes, M., Nuevo, F.A., Landi, E., Frazin, R.A.: 2017, Energy Input Flux in the Global Quiet-Sun Corona. *Astrophys. J.* **843**, 70. DOI. ADS.
- Nuevo, F.A., Huang, Z., Frazin, R., Manchester, i. Ward B., Jin, M., Vázquez, A.M.: 2013, Evolution of the Global Temperature Structure of the Solar Corona during the Minimum between Solar Cycles 23 and 24. *Astrophys. J.* **773**(1), 9. DOI. ADS.
- Nuevo, F.A., Vázquez, A.M., Landi, E., Frazin, R.: 2015, Multimodal Differential Emission Measure in the Solar Corona. *Astrophys. J.* **811**(2), 128. DOI. ADS.
- Press, W.H., Teukolsky, S.A., Vetterling, W.T., Flannery, B.P.: 2002, *Numerical recipes in C++ : the art of scientific computing*. ADS.
- Schiff, A.J., Cranmer, S.R.: 2016, Explaining Inverted-temperature Loops in the Quiet Solar Corona with Magnetohydrodynamic Wave-mode Conversion. *Astrophys. J.* **831**(1), 10. DOI. ADS.
- Serio, S., Peres, G., Vaiana, G.S., Golub, L., Rosner, R.: 1981, Closed coronal structures. II - Generalized hydrostatic model. *Astrophys. J.* **243**, 288. DOI. ADS.
- Spitzer, L.: 1962, *Physics of Fully Ionized Gases*. ADS.