

Magnetorotational instability: a review

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(Received 28 May 2013)

We review current research on magnetorotational instability (MRI) as a mechanism for momentum transportation in accretion disks.

1. Introduction

Background info

What is MRI? Accretion disks? Maybe provide references to explain relevance of MRI to birth of stars, geophysics, etc.

Also outline what the report will cover (probably as the last paragraph).

2. Physical Explanation

Physical explanation of magnetohydrodynamic (MHD) equations, concept of "frozen-in-field", analogy of magnetic and string tension, main idea of MRI

3. Theoretical Work

Acheson & Hide (1973) and Knobloch (1992) showed the linear stability analysis of rotating magneto-fluid bounded in coaxial cylinder. For the case of gaseous astrophysical disks in unbounded geometry was shown in Balbus & Hawley (1991). The condition for stability was shown to be radially increasing angular velocity profile. We present a linear stability analysis of MRI for bounded case following Acheson (1972), Acheson (1973), Knobloch (1992) and Julien & Knobloch (2010).

3.1. Governing equation

The wave dispersion equation of a cylindrical magneto-fluid can be obtained from the magnetohydrodynamic (MHD) equations. Assuming the fluid is inviscid and perfectly conducting, the ideal MHD equations are

$$\rho \left(\frac{\partial \mathbf{u}}{\partial t} + \mathbf{u} \cdot \nabla \mathbf{u} \right) = -\nabla p + \mathbf{J} \times \mathbf{B} \quad (3.1)$$

$$\frac{\partial \rho}{\partial t} + \nabla \cdot (\rho \mathbf{u}) = 0 \quad (3.2)$$

$$\mathbf{E} + \mathbf{u} \times \mathbf{B} = 0 \quad (3.3)$$

$$\nabla \times \mathbf{E} = -\frac{\partial \mathbf{B}}{\partial t} \quad (3.4)$$

$$\nabla \times \mathbf{B} = \mu \mathbf{J}. \quad (3.5)$$

To investigate MRI, one can consider homogeneous incompressible fluid rotating with angular velocity $\Omega(r) = \frac{V(r)}{r}$ in externally imposed magnetic fields $\mathbf{B}_0 = [0, B_\phi(r), B_z(r)]$.

Combining electromagnetic equations with momentum relation, we get the appropriate MHD equations,

$$\frac{\partial \mathbf{u}}{\partial t} + (\mathbf{u} \cdot \nabla) \mathbf{u} = -\frac{1}{\rho} \nabla \left(P + \frac{\mathbf{B}^2}{2\mu} \right) + \frac{1}{\mu\rho} (\mathbf{B} \cdot \nabla) \mathbf{B} \quad (3.6)$$

$$\frac{\partial \mathbf{B}}{\partial t} + (\mathbf{u} \cdot \nabla) \mathbf{B} = (\mathbf{B} \cdot \nabla) \mathbf{u} \quad (3.7)$$

$$\nabla \cdot \mathbf{u} = 0 \quad (3.8)$$

$$\nabla \cdot \mathbf{B} = 0 \quad (3.9)$$

3.2. Linear perturbation equation and eigenvalue problem

We can get linearized equations by perturbing the basic state by small amount of \mathbf{u}_1 and \mathbf{b}_1 for velocity and magnetic fields. The axisymmetric perturbation is assumed to have the form

$$\psi = \Re \left[\hat{\psi}(r) e^{i(kz - \omega t)} \right]. \quad (3.10)$$

According to Acheson (1972), the normal mode equations are

$$-i\omega \hat{b}_r = ikB_z \hat{u}_r \quad (3.11)$$

$$-i\omega \hat{b}_\phi = -\frac{d}{dr}(\hat{u}_r B_\phi) + ik(\hat{u}_\phi B_z - \hat{u}_z B_\phi) \quad (3.12)$$

$$-i\omega \hat{b}_z = -\frac{1}{r} \frac{d}{dr}(r \hat{u}_r B_z) \quad (3.13)$$

$$ik \hat{u}_z = -\frac{1}{r} \frac{d}{dr}(r \hat{u}_r) \quad (3.14)$$

$$i(\omega^2 - k^2 V_{Az}^2) \hat{u}_\phi = \frac{\hat{u}_r}{r} (2\Omega r \omega + 2k V_{A\phi} V_{Az}) \quad (3.15)$$

where $V_{A\phi} = \frac{B_\phi(r)}{\sqrt{\mu\rho}}$, $V_{Az} = \frac{B_z(r)}{\sqrt{\mu\rho}}$ are Alfvén speeds for associated external magnetic field components.

Solving the normal mode equation set for radial velocity perturbation $\hat{u}_r = u$, Acheson (1973) obtained following eigenvalue problem,

$$\begin{aligned} \frac{d}{dr} \left[(\omega^2 - k^2 V_{Az}^2) \left(\frac{du}{dr} + \frac{u}{r} \right) \right] - k^2 \left[\omega^2 - k^2 V_{Az}^2 + r \frac{d}{dr} \left(\frac{V_{A\phi}^2}{r^2} - \frac{V^2}{r^2} \right) \right] u \\ = -\frac{4k^2}{r^2} \frac{(k V_{A\phi} V_{Az} + \omega V)^2}{(\omega^2 - k^2 V_{Az}^2)} u. \end{aligned} \quad (3.16)$$

3.3. Stability criterion

3.3.1. Standard magnetorotational instability

Consider a standard MRI of radially bounded coaxial fluid cylinder with externally imposed axial magnetic field but without axial current flowing. Therefore, we have $V_{Az} = \frac{B_z}{\sqrt{\rho\mu}} = \text{const} \neq 0$ and $V_{A\phi} = \frac{B_\phi}{\sqrt{\rho\mu}} = 0$. Considering boundary condition $u(r_1) = u(r_2) = 0$, we multiply the eigenvalue equation by complex conjugate of u and integrate over radial coordinate,

$$(\omega^2 - k^2 V_{Az}^2)^2 = \frac{k^2}{D} \int_{r_1}^{r_2} \left[\frac{\omega^2}{r^2} \frac{d}{dr} r^2 V^2 - r^2 k^2 V_{Az}^2 \frac{d}{dr} \left(\frac{V^2}{r^2} \right) \right] |u|^2 dr \quad (3.17)$$

where

$$D \equiv \int_{r_1}^{r_2} \left(r \left| \frac{du}{dr} \right|^2 + \frac{|u|^2}{r} + k^2 r |u|^2 \right) dr > 0 \quad (3.18)$$

According to [Chandrasekhar \(1960\)](#), ω^2 must be real. We get stable modes with $\omega^2 > 0$ and unstable modes with $\omega^2 < 0$. If the angular velocity increases radially outward, $\frac{d}{dr} \left(\frac{V^2}{r^2} \right) > 0$, the system is stable because ω^2 is bounded from below by positive number,

$$\omega^2 > \frac{r^2 k^2 V_{Az}^2 \frac{d}{dr} \left(\frac{V^2}{r^2} \right)}{4 \frac{V^2}{r} + r^2 \frac{d}{dr} \left(\frac{V^2}{r^2} \right)} > 0. \quad (3.19)$$

If we have radially decreasing angular velocity profile, $\frac{d}{dr} \left(\frac{V^2}{r^2} \right) < 0$, somewhere $r_1 < r < r_2$, then ω^2 may have negative solution which makes the system unstable.

3.3.2. Helical magnetorotational instability

When the external nonzero magnetic fields in axial and azimuthal directions are considered, it was found that the eigenvalue equation can be written as

$$\frac{d}{dr} r \frac{du}{dr} - \frac{u}{r} - k^2 r u = \frac{k^2}{(\omega^2 - k^2 V_{Az}^2)^2} \left[r^2 \frac{d}{dr} \left(\frac{V_{A\phi}^2 - V^2}{r^2} \right) (\omega^2 - k^2 V_{Az}^2) - \frac{4}{r} (k V_{A\phi} V_{Az} - \omega V)^2 \right] u \quad (3.20)$$

According to [Knobloch \(1992\)](#) and [Julien & Knobloch \(2010\)](#), the exponentially growing mode $\omega = -i\lambda$, $\lambda > 0$ is possible when the eigenvalue relation has following form

$$(\lambda^2 + k^2 V_{Az}^2)^2 = \frac{k^2}{D} \int_{r_1}^{r_2} \left[r^2 \frac{d}{dr} \left(\frac{V_{A\phi}^2 - V^2}{r^2} \right) (\lambda^2 + k^2 V_{Az}^2) + \frac{4}{r} (k V_{A\phi} V_{Az} - i\lambda V)^2 \right] |u|^2 dr. \quad (3.21)$$

Considering the imaginary part of the equation, we have

$$\int_{r_1}^{r_2} \frac{1}{r} V_{\phi} V |u|^2 dr = 0 \quad (3.22)$$

[Knobloch \(1992\)](#) shows exponentially growing instability is only possible when $V_{A\phi}$ or V changes sign somewhere in $r_1 < r < r_2$.

4. Laboratory Experiments

Discuss on experimental work that has been done related to this topic.

Try to categorize experiments so we can group the discussion. Also, the main focus is to explain what has been done to attempt to recreate MRI in the lab.

NOTE: most literature on MRI experiments seem to be about either 1) axial-only magnetic fields (standard MRI) or 2) axial and azimuthal magnetic fields (helical MRI).

5. Numerical Work

MRI: occurs in a Rayleigh-Stable regime whenever a weak poloidal magnetic field is present.

Accretion can only occur with a efficient mechanism for outward transport of angular

momentum. There are many suggested processes that can lead to efficient accretion but none leads to substantial outward angular momentum transport. Before using any model, there is a need to understand the angular momentum transport characteristics so that the appropriate equations can be used. The flow could be turbulent but there is no clear reason why there would be turbulence since there is no known instability that will encourage turbulent flow.

Numerical simulations are employed to try to explain the anomalous viscosity and magnetic field generation. Viscous stresses is insufficient for explaining the outward transport of angular momentum so there is a need for some other mechanism. Purely hydrodynamic flow will not generate that as will be shown by HGB and a weak magnetic field is needed. A strong magnetic field will suppress MRI as shown by [Liu \(2008\)](#). From linear stability analysis of the exact nonlinear streaming solution done by [Goodman \(1994\)](#), we know the magnetic field perturbations grows exponentially, indicating that the solution is unstable. 3D numerical simulation will help to confirm this analysis.

Direct numerical simulations (DNS) are done to solve the magnetohydrodynamic equations for: compressible and incompressible; ideal and non-ideal; 2D and 3D; axisymmetric and non-axisymmetric; field strength; resolution; computational domain size; incorporation of coriolis; and tidal forcing.

With regards to 2D vs 3D simulations, a good comparison was done by Hawley et al HGB 1995 where they did a local shearing box model which incorporated coriolis and tidal forcing but neglects background gradients in pressure and density (not important unless radial oscillations are significant). Periodic boundary conditions were used with an added shearing component to the radial direction. "Method of characteristics-constrained transport" algorithm as explained in [Hawley \(1995\)](#) was implemented.

The comparison of the 3D simulations results from HGB with the 2D results from Balbus et al BGH 1994 shows significant difference. The 2D results have a general channel/streaming flow solution whereas the 3D results may go through the channel flow but eventually evolves further into a turbulent flow with some not going through the channel flow at all. Note: vertical component of gravity, hence buoyancy, was ignored and since their simulations showed significant density contrast, the shearing box model is not self-consistent. Purely hydrodynamic turbulence did not give significant angular momentum transport needed even when the initial conditions had strong turbulence structures. Their 3D results suggest that accretion disks are well described by Euler's, ignoring the viscous terms from Navier Stokes.

Squire's theorem says that for each unstable 3D disturbances, there are corresponding 2D disturbance that are more unstable. How does that apply here?

Non-axisymmetric [Fleming \(2000\)](#) - suggest saturation does occur even when Elsasser number $\Delta \geq 1$ if non-axisymmetric disturbances are allowed to evolve.

If Alfvén speed is slower than sound speed, equations can be simplified to be incompressible since it is not fundamental to the instability [Balbus & Hawley \(1991\)](#) ...

Incompressible shearing box : [Lesur \(2007\)](#)

[Liu \(2008\)](#) (also Princeton MRI experiments) simulated nonlinear development of MRI in a non-ideal magnetohydrodynamic Taylor-Couette flow (mimicking an on-going experiment) using ZEUS-MP 2.0 code ([Hayes \(2006\)](#)), which is a time-explicit, compressible, astrophysical ideal MHD parallel 3D code with added viscosity and resistivity for axisymmetric flows in cylindrical coordinates. He shows that the saturation of MRI causes a inflowing 'jet' feature which is opposite to the usual Ekman circulation and enhances angular momentum transport radially outward, which is what is needed, agreeing with HGB. Note: Stable MRI regime ($Re \leq 1600$) enhances vertical angular momentum trans-

port while in unstable MRI regime ($Re \geq 3200$), MRI kicks in, resulting in more radial angular momentum transport as compared with vertical.

The use of the shearing box model has its limitations as pointed out by HGB as they do not permit any dynamics involving the background shear, which is taken as imposed and constant in space and time. [Regev \(2008\)](#). With regards to resolution which is tied in with the model used, the efficiency of angular momentum transport appears to decrease with improved resolution as shown by [Fomang \(2007\)](#), suggesting that transport rates estimated on basis of ideal MHD are affected by grid-scale dissipation and overestimate the efficiency of angular momentum extraction by MRI. [Kopyla \(2008\)](#) also worked on resolution dependence of α , the Shakura-Sunyaev number and the possible decline of the angular momentum transport rate with decreasing Pm.

Sawai et al 2013 [Sawai \(2013\)](#) simulated a global 3D MRI axisymmetric, ideal model. They did not find any channel flow solutions as HGB did. Increasing spatial resolution contributed to approximate convergence in the exponential growth rate but not that of the saturation of magnetic field due to large numerical diffusion.

[Kirillov et al. \(2012\)](#) presents a unifying description of the helical and azimuthal versions of MRI, and they also identify the universal character of the 'Liu' limit $2(1-2) \approx -0.8284$ for the critical Rossby number. (From this universal characteristics, they are led to the prediction that the instability will be governed by a mode with an azimuthal wavenumber that is proportional to the ratio of axial to azimuthal applied magnetic field, when this ratio becomes large and the Rossby number is close to the Liu limit)

[Miller \(1999\)](#) worked on vertically stratified disks using a shearing box model and found that turbulent magnetised disk can produce a magnetised corona in laminar flow through MRI.

The influence of viscosity and electrical resistivity on the MRI ([Pessah \(2008\)](#))

From [Julien & Knobloch \(2010\)](#): Reduced equations- Case A ($\delta = \epsilon, \Lambda = O(1)$)

- a) Single Mode theory
- b) Stress-free boundary conditions
- c) Random initial conditions

Reduced equations- Case B ($\epsilon = O(\delta), \Lambda \gg 1$)

- a) Single Mode Solutions
- b) Multi-mode Solutions
- c) Random initial Perturbations
- d) Dissipation and Saturation

Note: Accretional ejection instability (AEI), as suggested by [Caunt \(2011\)](#), gives the same angular momentum transport needed as compared to the global MRI but the instability requires a fairly strong magnetic field. So need an external source, possibly from MRI.

6. Conclusion

Concluding remarks

REFERENCES

- ACHESON, D. J. 1972 On the hydromagnetic stability of a rotating fluid annulus. *Journal of Fluid Mechanics* **52**, 529–541.

- ACHESON, D. J. 1973 Hydromagnetic wavelike instabilities in a rapidly rotating stratified fluid. *Journal of Fluid Mechanics* **61**, 609–624.
- ACHESON, D. J. & HIDE, R. 1973 Hydromagnetics of rotating fluids. *Reports on Progress in Physics* **36** (2), 159.
- BALBUS, S. A. & HAWLEY, J. F. 1991 A powerful local shear instability in weakly magnetized disks. i - linear analysis. ii - nonlinear evolution. *Astrophysical Journal* **376**, 214–233.
- CAUNT 2011 .
- CHANDRASEKHAR, S. 1960 The stability of non-dissipative couette flow in hydromagnetics. *Proceedings of the National Academy of Science* **46**, 253–257.
- FLEMING 2000 .
- FOMANG 2007 .
- GOODMAN 1994 .
- HAWLEY 1995 .
- HAYES 2006 .
- JULIEN, K. & KNOBLOCH, E. 2010 Magnetorotational instability: recent developments. *Philosophical Transactions of the Royal Society A* **368**, 1607–1633.
- KAPYLA 2008 .
- KIRILLOV, OLEG N., STEFANI, FRANK & FUKUMOTO, YASUhide 2012 A unifying picture of helical and azimuthal magnetorotational instability, and the universal significance of the liu limit. *The Astrophysical Journal* **756** (1), 83.
- KNOBLOCH, E. 1992 On the stability of magnetized accretion discs. *Monthly Notices of the Royal Astronomical Society* **255**, 25P–28P.
- LESUR 2007 .
- LIU, WEI 2008 Numerical study of the magnetorotational instability in princeton mri experiment. *The Astrophysical Journal* **684** (1), 515.
- MILLER 1999 .
- PESSAH 2008 .
- REGEV 2008 .
- SAWAI 2013 .