

Temporal and Spatial Turbulent Spectra of an Experimental MHD Wind-tunnel Plasma and an Observation of Variance Anisotropy

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The nature of MHD turbulence in an experimental device is analyzed through both temporal and spatial magnetic fluctuation spectra. A magnetically turbulent plasma is produced in the MHD wind-tunnel configuration of the Swarthmore Spheromak Experiment. The power of magnetic fluctuations are projected into directions perpendicular and parallel to a local mean field; the ratio of these quantities shows the presence of variance anisotropy which varies as a function of frequency. Comparison amongst magnetic, velocity, and density spectra are also made, demonstrating that the energy of the turbulence observed is primarily seeded by magnetic fields created during plasma production. Direct spatial spectra are constructed using multi-point diagnostics and are used to compare to frequency spectra converted to spatial scales using the Taylor Hypothesis. All turbulence results compare fairly well to a Hall-MHD simulation of the experiment. Evidence for the observation of dissipation due to ion inertial length scale physics is also discussed as well as the role laboratory experiment can play in understanding turbulence typically studied in space settings such as the solar wind.

I. INTRODUCTION

The solar wind turbulence community has seen a flurry of results in recent years, as diagnostic capabilities on spacecraft have steadily improved, especially in the temporal resolution which has allowed for investigation into ion, sub-ion and electron scale turbulence measurements. In addition, newer analysis techniques have been developed to better tap the details of the rich turbulent environment that is the heliosphere.

Analysis of timeseries fluctuations in the guise of spectrum analysis has served as the foundation for turbulence research in the solar wind^{1,2}. Improving diagnostics has led to ever greater precision in measurements of spectral indices for both magnetic and velocity fluctuation spectra³, as well as increasing resolution of scale, allowing for measurements of spectra of sub-ion and electron scale fluctuations⁴⁻⁶. However, many outstanding questions regarding the nature of this turbulence, including both injection and dissipation mechanisms, have prompted research beyond typical spectral analysis. Two such avenues have been the exploration of the anisotropic nature of the turbulence and comparison of spectra between different observables (i.e. magnetic versus density fluctuations).

Theoretical treatments of magnetized plasma turbulence have almost universally predicted anisotropy to develop based on directions perpendicular and parallel to a mean field vector⁷⁻¹¹; many forms of this anisotropy have been predicted and observed^{12,13,15,38} (a good overview of the various types is given in Horbury 2012)¹⁶. Since the solar wind speeds are typically super-Alfvenic, anisotropy studies of solar wind turbulence generally reference the velocity vector of the bulk flow when defining perpendicular versus parallel fluctuations—typically resulting in a wavenumber anisotropy measurement. This paper focuses on variance anisotropy—the observation of greater

magnetic fluctuation power in components perpendicular to the mean magnetic field versus parallel^{17,18}. From the experimentalist point of view, this is the most straightforward quantity to report and a reasonable place to begin.

Comparison amongst fluctuation spectra (typically magnetic, flow, and density fluctuations) also provide insight into the nature of the turbulence. For example, relative partition of fluctuation energy between magnetic and kinetic can be useful in making determinations of turbulent properties³ or comparison of the spectral indices of magnetic and density spectra suggest an increase in the compressibility of a plasma (the so-called density bulge^{19,20}). Comparison of velocity and magnetic spectra at injection scales can also be used to develop hypotheses on the generating and dissipation mechanisms of the turbulence³⁶.

As results from space become more detailed, comparison to simulation and laboratory experiment become more useful in order to better develop the theory as well as inform future space missions. A long gap in laboratory research has occurred since the earliest observations of magnetic turbulence and anisotropy²² as turbulence work in experiments gravitated toward electrostatic turbulence and transport in tokamaks, but more recently, much progress has been made to develop laboratory experiment which can inform space plasma turbulence research²³ as well as begin to make magnetic turbulence measurements²⁴. Experiments conducted in the MHD wind-tunnel configuration of the Swarthmore Spheromak Experiment have shown the ability to produce and analyze MHD turbulence using many of the same methods used in the space plasma community, and has produced turbulence measurements which are comparable to *in-situ* results²⁵. Moreover, using varying amounts of helicity injection, some characteristics of the turbulence can be modified²⁶.

A full spectral analysis of fluctuations in the SSX

has been conducted, including magnetic field, density and Mach flow fluctuations. Using a wavelet transform method to decompose the timeseries signal, the magnetic field fluctuation spectra can be broken into portions that are parallel or perpendicular to the local magnetic field vector. Analysis of these portions show that parallel fluctuation power decreases slightly faster than that for perpendicular fluctuations generating a separation in power as a function of increasing frequency. This effect in frequency space is very close to the observation of variance anisotropy seen in wavenumber space in solar wind turbulence²⁷. The ratio is shown to grow from nearly isotropic levels in the energy injection scale to a maximum \perp / \parallel ratio of 3 at frequencies of about 1MHz. Beyond this point, the ratio begins to return to isotropic levels. Though the axial flow speeds of the plasma in the wind-tunnel sub-sonic ($M_z < 1$), if a Taylor hypothesis were still invoked this peak in the ratio would correspond to the ion inertial scale length.

Since density and Mach fluctuations timeseries are also recorded, comparisons amongst these spectra are also reported. Comparison of Mach number fluctuation spectra (as a proxy for flow) to magnetic spectra in the lowest frequencies suggests that energy injection into the turbulent spectra is primarily magnetic. This result is consistent with the nature of the spheromak generation process. Also the indices of density and magnetic spectra appear to diverge at a scale consistent with the ion inertial length and also where anisotropy is found to decrease. Both results would be consistent with an increase in compressibility of the plasma due to ion scale physics. However, firm conclusions are not made as noise issues in the density and flow diagnostics at that scale may be a factor.

Finally, direct wavenumber spectra are made using a multi-tipped probe and is compared to time-domain spectra transformed to spatial scales using the Taylor Hypothesis. This diagnostic highlights an advantage of laboratory experiments since actual spatial measurements are extremely difficult *in-situ* though progress is being made using multiple spacecraft²⁸. The wavenumber spectra tend to be slightly shallower than the time-domain counterpart. Comparison of the wavenumber spectra with respect to flow gives a hint at an observation of wavenumber anisotropy, but the results do not provide as clear evidence as for the variance anisotropy analysis.

All the temporal and spatial spectra analyzes are compared to a Hall-MHD simulation generated using the HiFi framework²⁵. Spectra are generated from synthetic diagnostic timeseries of magnetic fields, flow, and density. Results compare favorably to the experiment including the observation of variance anisotropy.

The organization of this paper is as follows: First, a brief description of the plasma laboratory is given. Then, an overview of the analysis techniques used to determine perpendicular and parallel fluctuations is provided; details of the method and a second method used to verify its validity are described in the appendix. The main re-

sults of this paper are provided in Sections IV to VI: Section IV presents the variance anisotropy results, Section V compares magnetic fluctuation spectra with velocity and density and Section VI shows the results of the direct wavenumber spectra analysis. Section VII presents the comparison of simulation with experimental results. Section VIII discusses the results in light of its relationship to possible ion inertial length scale dissipation. Concluding remarks are given in Section IX. The appendices give details of the variance anisotropy analysis as well as a discussion of Wavelet versus FFT transforms and local versus global magnetic fields.

II. EXPERIMENT

A turbulence cascade process in a laboratory experiment is naturally going to have a different origin than a space physics process; however, it is assumed that processes after the energy injection state (i.e. energy transfer in the inertial range and dissipation) will be similar enough so that exploration of the physics behind them in the laboratory can be beneficial to an overall understanding.

The energy for the turbulence found in the MHD wind-tunnel in the Swarthmore Spheromak Experiment originates in the plasma production process. As diagrammed in Figure 1, a plasma gun configuration sits on one end of a 15.5cm diameter, 86cm long cylindrical copper column which constitutes the MHD wind-tunnel. The gun consists of a tungsten-coated 4cm diameter inner electrode placed concentrically within the copper cylinder which serves as an outer electrode. An axial aligned wire coil surrounds both electrodes and current is supplied to the coil to produce a known amount of magnetic flux—between 0 and 1.5mWb—axially through the inner electrode: this flux is referred to as stuffing flux, Φ . A 1mF capacitor bank, charged to 4kV is discharged across the electrodes; this voltage fully ionizes a small volume of hydrogen gas puffed in just before the discharge. Radial currents through this newly produced plasma push the plasma down the column and into the fringe magnetic fields which tend to resist this push and stuff the progress of the plasma (hence the term, stuffing flux). Given enough current, and thus large enough $J \times B$ force, the plasma distends the stuffing fields until they break off, forming a self-contained magnetic field structure called a spheromak^{29,31}. This structure is visualized in Figure 1(a) using Hall-MHD simulation generated field lines (in blue). Since the spheromak has both poloidal and toroidal magnetic fields, the relative ratio of field strength between these two directions is quantified by the magnetic helicity, defined as,

$$K_B = \int A \cdot B dV \quad (1)$$

where A is the vector potential and dV is the volume element. This helicity can be thought of as the internal

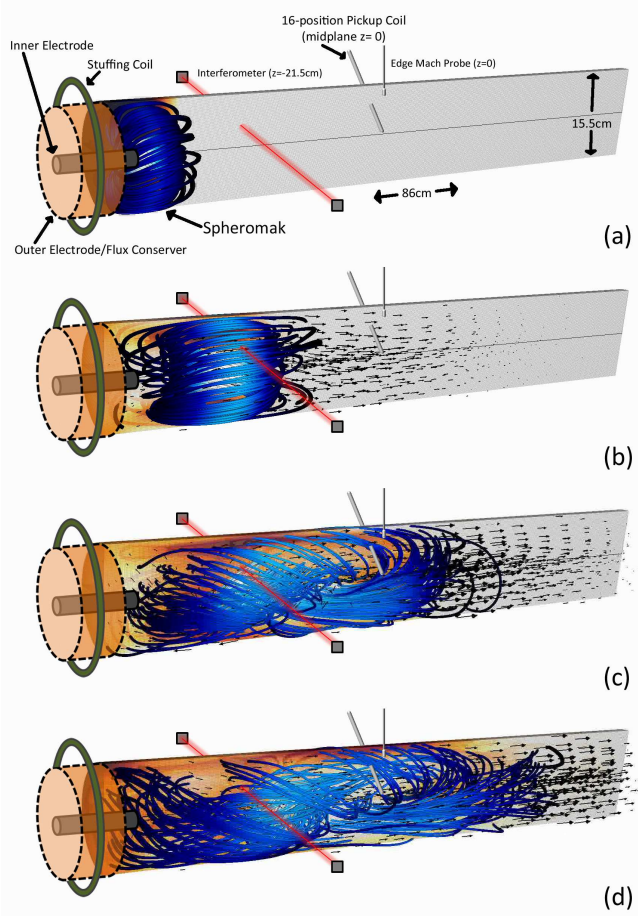


FIG. 1. A diagram of the MHD Wind-tunnel inside the SSX chamber. A plasma gun is located on one end and consists of two electrodes and a stuffing coil. Magnetic pickup coils are located at 16 radial locations at the midplane and a Mach probe measures edge flow at a different angular position at the midplane. The position of the density HeNe interferometer chord is also indicated. Panels (a)-(d) demonstrate the plasma evolution once a spheromak has been formed at the end the plasma gun(a), and shows the tilt instability(b), twisting under helicity conservation(c) and measurement of turbulent fluctuations(d). The blue lines show simulation generated magnetic field lines and the arrows indicate plasma flow magnitude and direction.

twist of the magnetic field line structure. Previous work on SSX has shown that magnetic helicity of the plasma scales approximately linearly with the amount of stuffing flux applied to the gun²⁶

Figure 1(a)-(d) illustrates the experimental procedure. The generalized turbulence cascade begins with this compact magnetic structure (Figure 1(a)). Inside the wind-tunnel, the magnetic structure is energetically unstable³¹—the structure will begin to tilt over and expand into the remainder of the wind-tunnel (Figure 1(b)). Because the column is copper, and thus flux conserving, the magnetic helicity is conserved unlike the magnetic energy⁷; thus, the structure also begins to twist as it tilts over as seen in Figure 1(c). The free energy released in the fall-over materializes as fluctuations in the field, generating the turbulent cascade. The turbulent fluctuations are most prominent in figure 1(d). In the actual experiment, the gun typically injects more than a single, self-contained structure so while an initial structure is decaying, more compact field energy is being injected²⁹. This allows for a time frame of stationary fluctuations that is used in the turbulence analysis.

The turbulence data is extracted from magnetic, density and flow measurements during this stationary period. Magnetic fluctuations are recorded using 3mm diameter, single loop pick-up coils located at 16 radial locations along the radius of the midplane of the column as indicated in Figure 1(a). Each radial position has three orthogonal loops oriented along the axial, radial and azimuthal directions of the column. A 64MHz, 14-bit DTaq digitizer records $dB/dt = \dot{B}$ timeseries data which is converted into magnetic fluctuation data in frequency space (as discussed in Section III). Line-integrated density data is measured with a HeNe interferometer located 21.5cm off of the midplane and flow fluctuations are estimated from a Mach probe located on the edge of the copper column at the midplane (as indicated in Figure 1(a)). Spectra of $M_z(t)$ are directly reported as a proxy for $V_z(t)$, since $\tilde{M} \sim \tilde{V}/C_s$ where $C_s = (T_e/m_i)^{1/2}$ and is constant given a constant measured value of $T_e = 10eV$ ^{26,30}. Bulk flow of the plasma is estimated with time-of-flight measurements between the density signal at $z=-21.5cm$ and the magnetic signal at the midplane, $z=0$. The plasma is also generated with a set amount of magnetic helicity which is governed by initial conditions of the plasma gun source—namely, amount of flux generated in the gun core (the stuffing flux). The helicity can be scanned²⁶, but in this work focus is primarily on two states: a state with non-zero helicity, $K_B \neq 0$, generated by 1.0mWb of flux in the gun core and a state with no injected helicity, $K_B = 0$. Table I indicates typical plasma parameter values for these two states.

Each discharge of the plasma gun constitutes an experimental shot and typically lasts for about 120 μs from formation of the magnetic fields to their eventual resistive dissipation. The turbulence data reported here is extracted from a time range of 40 to 60 μs . This is the epoch during each shot where the fluctuations are most

TABLE I. MHD wind tunnel plasma parameters during the equilibrium epoch for the present configuration of SSX for non-zero helicity ($K_B \neq 0$) and for zero helicity ($K_B = 0$) states. The table has separate sections for directly measured parameters and for quantities computed from these values.

Parameter	$K_B \neq 0$ (1.0mWb)	$K_B = 0$ (0.0mWb)
<i>Measured</i>		
$\langle B \rangle [kG]$	5.283	0.747
$\langle n \rangle \times 10^{15} [cm^{-3}]$	1.39	2.84
$\langle T_i \rangle [eV]$	23	17
$V_{bulk} [km/s]$	20	20
<i>Computed</i>		
β	0.07	5.5
$V_a [km/s]$	309	30
$C_s [km/s]$	31	31
$\rho_i [cm]$	0.09	0.56
$\delta_i [cm]$	0.61	0.43
$\lambda_m f p^i [cm]$	0.16	0.05
$f_{ci} [MHz]$	8	1.1
$\nu_i [MHz]$	6	19
$f_{\delta i} [MHz]$	3	4.3
$f_{\rho i} [MHz]$	20	3.3

stationary; energy at the probe location is balanced between injection energy from the gun and loss through advection away from the probe and through possible dissipation mechanisms (including resistive decay of the currents).

An ensemble average for each helicity state is constructed from 40 shots.

III. ANALYSIS TECHNIQUES

The bulk of the data of this paper is analyzed in frequency space using either a Wavelet transform method⁴⁰ or a traditional Fast-Fourier Transform (FFT). Time-series data is transformed using a sixth-order Morlet mother wavelet with 256 logarithm scale steps per octave. For magnetic fluctuations, the wavelet transform is applied directly to the $\vec{B}(t)$ data from the pickup coils, separately for each orthogonal component, x . Each shot contains 8192 timesteps and the transform is applied to the entire shot to yield an $W_x(f, t)$. A known conversion factor is applied to convert the wavelet scales into Fourier frequencies⁴⁰. The transforms are divided through by the square of the frequency to convert $\dot{B}(f)$ to $B(f)$. Then, the transform is summed over the time range of interest, typically 40-60 μs , yielding,

$$B_x(f) = \frac{1}{f^2} \sum_{t=t_0}^{t_1} W_x(f, t) \quad (2)$$

. Since the entire shot is being used in the wavelet transform, the frequency range can extend beyond that which would be typical for an FFT given the particular time range under analysis. The entire frequency range is displayed in the following plots, but the focus of the analysis will be on the same frequency range accessible by

an FFT. Details on the comparison between the two approaches is given in Appendix B

A similar procedure is applied to the density timeseries from the interferometer and the Mach number timeseries from the Mach probe, but without the frequency scaling used for the magnetic data. For spatial spectra (using separate radial points on the magnetics probe) an FFT is used.

The method for conversion of the magnetic fluctuation spectra from $\vec{B} = (B_r, B_\theta, B_z)$ to $\vec{B} = (B_\parallel, B_\perp)$ for the variance anisotropy analysis is detailed in Appendix A, but utilizes the temporal resolution afforded by the Wavelet transformation.

IV. VARIANCE ANISOTROPY

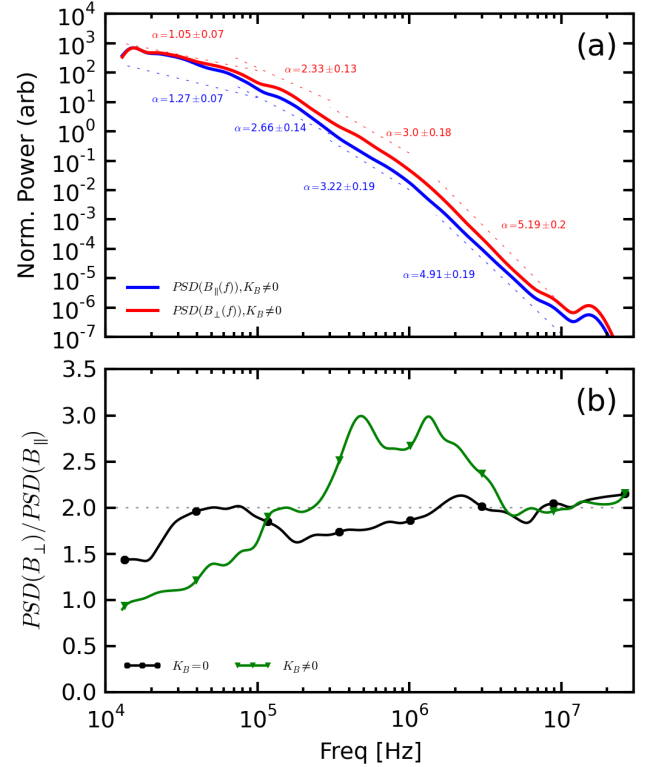


FIG. 2. (a) Magnetic power spectra of the non-zero helicity state portioned into perpendicular (red, higher) and parallel (blue, lower) components with respect to the local magnetic field vector direction. Power law fits and error are indicated for four sections of the spectra. Values and fit regions are indicated in Table II. (b) Ratio of the perpendicular over parallel power spectra from (a) as a function of frequency, plotted on a log-linear scale. The green curve (with triangles) shows the $K_B \neq 0$ state while black (with circles) shows the zero helicity state. An isotropic ratio is indicated by the horizontal dotted gray line.

The magnetic field fluctuation spectra perpendicular and parallel to the local magnetic field vector is shown

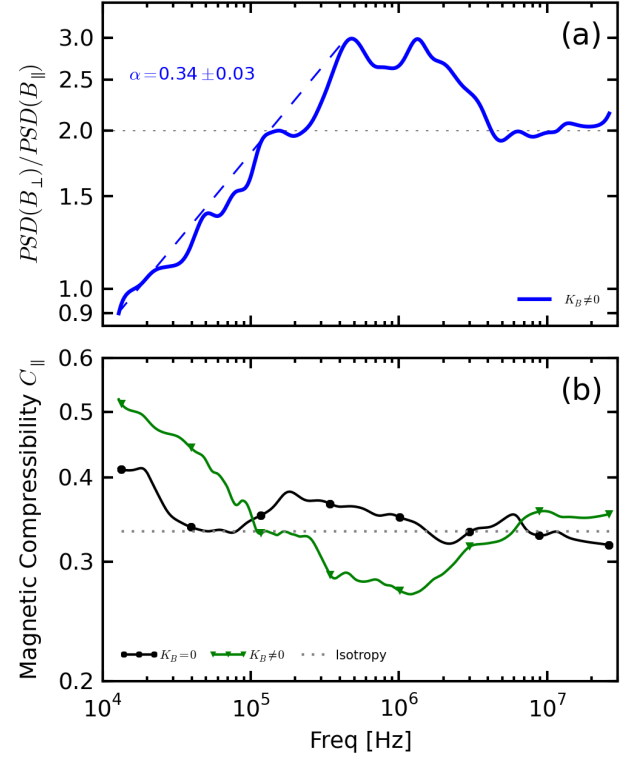
in figure 2(a) averaged over 40 shots and the inner four probe tips (a spatial range of 1.5cm). Like previously reported magnetic spectra²⁵, both perpendicular and parallel curves exhibit power-law like behavior for most frequencies between 10kHz and 10MHz. Power-law fits to $f^{-\alpha}$ are made to various sections of the curve using a Maximum Likelihood Estimation method³³; the short fits show that the scaling appears to gradually change from spectral indices of about 1 in the injection (or outer) range and from about 2.5 to 5 in the remaining frequency sections. As has been noted before, with the exception of the injection range slope, the inertial and possibly dissipation range spectral indices are steeper than observed in solar wind turbulence spectra.

The separation between perpendicular and parallel spectra, however, proceed in a manner similar to that found in space. Both perpendicular fluctuations (red) and parallel fluctuations (blue) begin at roughly the same magnitude in power. Beyond 20kHz, though, the parallel curve dips down slightly faster than the perpendicular and this trend continues up to about 500kHz. Then the separation begins to decrease and approaches a constant value from 5MHz to the Nyquist limit of 32MHz. This gradual change as a function of frequency is more clearly observed in figure 2(b) plotted in log-linear format, which provides the ratio of the two curves in figure 2(a). Since the fluctuations perpendicular to the B-field have two component directions while parallel fluctuations have only one, isotropy—equal fluctuation power in all three components—occurs when the ratio of perpendicular to parallel is two. Isotropy is indicated in figure 2(b) by the dashed gray line.

At the lowest frequencies, the balance of power actually tips toward parallel over perpendicular. As the frequency increases, the ratio approaches isotropy, then changes over to dominantly perpendicular power. The ratio continues to steadily increase reaching a peak of about 3 at about 500kHz. There is a brief dip before the ratio rises again at about 1.5MHz; at this point, the ratio begins to drop steadily, and approaches an isotropic level at about 5MHz.

The scale dependency of the $K_B \neq 0$ data can be compared to the zero-helicity state. The black curve shows the anisotropy ratio for the zero helicity state, where the field strength is about an order of magnitude less, and there is much less structure to the field. As might be expected for such a state, which has a larger β ^{18,34}, there appears to be very little anisotropy at any scale with values ranging close to $R=2$.

Fits to both the spectra and the ratio are shown in figures 2(a) and figure 3(a). In figure 2(a), fits are made to four sections of each curve. The spectral indices, error and range of fit for both parallel and perpendicular spectra are indicated in the figure and reproduced in table II. Overall, comparison of the spectral indices also reflect anisotropy: for each section fit, the parallel slope is consistently steeper than the perpendicular slope. This difference is also reflected in the fit shown in figure 3(a).



plot_wavelet_spec_anisotropy_forSpectrapaper.py

FIG. 3. (a) The same ratio of the green curve in Figure 2(b), but now plotted in a log-log format to highlight the power-law behavior of the ratio as a function of frequency. A fit between 10kHz and 500kHz is made with a spectral index of 0.34 indicated above the curve. (b) Magnetic compressibility, defined in Equation 3, for non-zero helicity (green, triangles) and zero helicity (black, circles) states as a function of frequency. An isotropic ratio is indicated by the horizontal dotted gray lines in both (a) and (b).

TABLE II. Indices from MLE power-law fits of magnetic fluctuation spectra in Figure 2.

Fit Range [MHz]	Direction	Index, α for $f^{-\alpha}$	Error
0.01-0.1	\parallel	1.27	0.07
	\perp	1.05	0.07
0.07-0.3	\parallel	2.66	0.14
	\perp	2.33	0.13
0.3-1.0	\parallel	3.22	0.19
	\perp	3.00	0.18
1.5-10	\parallel	4.91	0.19
	\perp	5.19	0.20

An index of $\alpha = 0.34$ indicates that the ratio scales approximately like $f^{1/3}$ for the region between 10kHz and 500kHz.

An alternate presentation of this ratio is shown in figure 3(b), where the ratio of perpendicular and parallel fluctuation power is cast in terms of magnetic compressibility²⁷,

$$C_{\parallel}(f) = \frac{1}{N} \sum_{j=1}^N \frac{B_{\parallel}(t_j, f)}{B_{\parallel}(t_j, f) + B_{\perp}(t_j, f)} \quad (3)$$

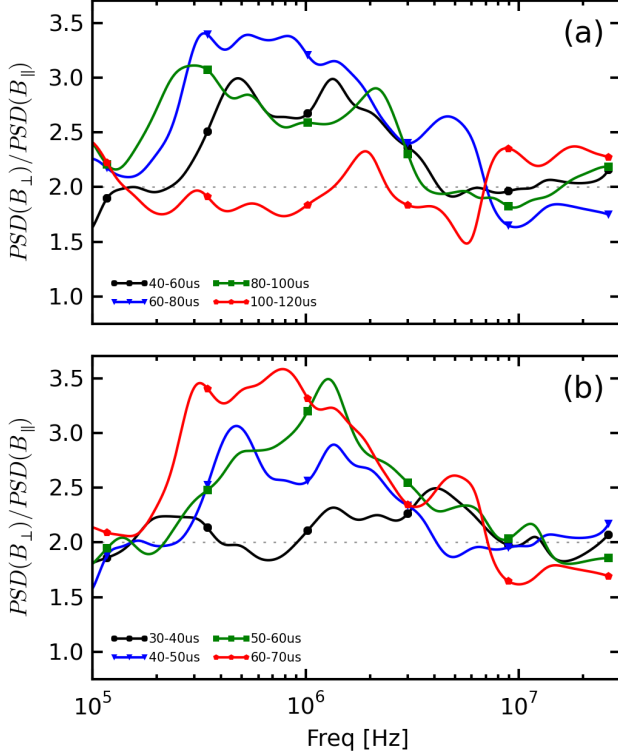


FIG. 4. (a) Anisotropy ratio for $K_B \neq 0$ data for four $20\mu s$ time ranges including and following the nominal analysis range of $40 - 60\mu s$. (b) Anisotropy ratio for $K_B \neq 0$ data for four $10\mu s$ time ranges just before, including, and just after the nominal analysis range. An isotropic ratio is indicated by the horizontal dotted gray lines.

which relates the anisotropy to the characteristic stiffness of the magnetic field structure. Note the ratio between parallel and total is taken before the Wavelet spectra are summed over the time range with this definition while the pure ratio shown in previous figure is taken after integration over time. This metric can be used to help unravel the nature of the fluctuations^{27,35}.

An evolution of the anisotropy over time is also observed. Figure 4(a) and (b) shows the change of the anisotropy ratio as a function of time ranges of $20\mu s$ intervals(a) and $10\mu s$ intervals(b). The black curve in figure 4(b) shows a period of time, $30-40\mu s$, just as the plasma is reaching the midplane probe. The curve remains near isotropic levels for most of the frequencies. As the time frame increases in $10\mu s$ intervals, the perpendicular power clearly increases in the 100kHz to 3MHz range. The ratio actually peaks highest in the time frame just beyond the main analysis period shown in figure 2(b). These trends perhaps demonstrate that the anisotropy increases as more time is allowed for the turbulence to evolve. Figure 4(a), shows that after energy is no longer being injected to maintain the turbulence, the anisotropy decreases. After $100\mu s$ (red curve), the magnetic field fluctuations have returned to an isotropic ratio.

V. FLOW AND DENSITY SPECTRA

For a turbulence cascade to develop, a system needs both energy injection and energy dissipation. The separation of spatial scale between injection and dissipation determines the size of the inertial range. While the actual injection mechanisms of the solar wind are not completely understood, there is evidence from the comparison of large scale magnetic field and velocity fluctuation data that velocity fluctuation energy is being tapped by magnetic fluctuations to sustain an injection-range like cascade for magnetic spectra³⁶. In the SSX plasma, however, the injection scale energy is primarily magnetic—the formation of the unstable spheromaks. This is borne out by similar comparisons of magnetic spectra and velocity spectra. Figure 5 shows magnetic spectra (a) and Mach number fluctuation spectra (b) for the two helicity states. As indicated in Table I, the non-zero and zero helicity states also have fairly different magnetization levels— 5kG vs 0.8kG , respectively.

Comparison of the blue and red curves in Figure 5(a) show that the larger mean magnetic fields of the non-zero helicity state are correlated with power in low frequency fluctuations while the lower magnetized state has effectively noise at these lower frequencies. Comparing these states with the respective velocity fluctuations seems to indicate that the additional magnetic energy injection at larger scales affects the velocity fluctuation power. For the low field/zero-helicity state, the velocity cascade is immediately steep as well as lower in energy overall. In the high field/non-zero helicity state, however, the velocity fluctuation scaling is shallower, but then has a breakpoint around the frequency where B-field fluctuation energy is about the same between high and low field states. Furthermore, the higher magnetic energy of the $K_B \neq 0$ correlates with higher overall flow fluctuation power of about an order of magnitude. These results suggest that the state with more injected magnetic energy delivers some of its energy to the velocity fluctuations, though more direct evidence of this connection is still being sought.

Figure 6 shows non-zero helicity state magnetic and velocity fluctuation spectra on the same relative scale with power-law fits indicated in the figure and in Table III. The velocity spectra scales like f^{-1} about to about 200kHz , a scaling which is typical for injection range turbulence. Meanwhile, the magnetic spectra is steeper, which implies that some of the magnetic energy may be going to drive flows, though the exact mechanism of this energy transfer is not known and its study is reserved for later work. The flattening of the velocity spectra could also be a signal of compressive effects³⁶.

Beyond 200kHz , the velocity fluctuations scaling steepens suggesting an inertial range scale. The B-field steepens further as well. Then, the curves match slopes for about a decade in frequency space. Beyond, 2MHz , the B-field scaling drops off significantly while the velocity spectra scaling actually slightly increases. This is pos-

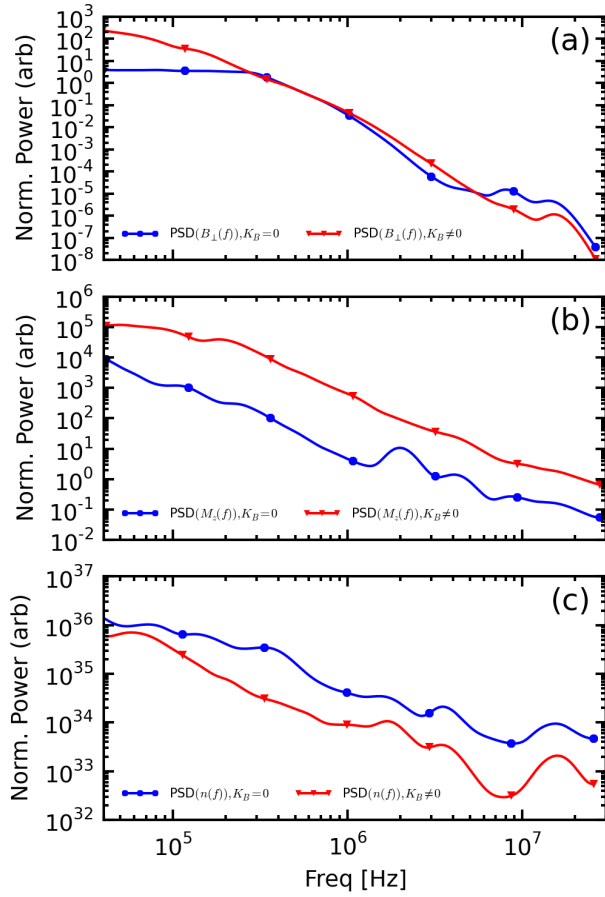


FIG. 5. Comparison of spectra between the non-zero helicity ($K_B \neq 0$, red triangles) and zero helicity ($K_B = 0$, blue circles) for (a) magnetic power spectra, (b) Mach number power spectra (as a proxy for velocity spectra) and (c) line-integrated density power spectra. The levels of the curves in this plot indicate the relative power between the quantities in the two different helicity states.

TABLE III. Indices from MLE power-law fits of magnetic and Mach number fluctuation spectra in Figure 6.

Fit Range [MHz]	Parameter	Index	Error
0.05-0.2	M_z	1.03	0.12
	B_\perp	2.03	0.13
0.3-1.0	M_z	2.70	0.17
	B_\perp	3.00	0.18
2.0-10	M_z	2.10	0.11
	B_\perp	5.17	0.21

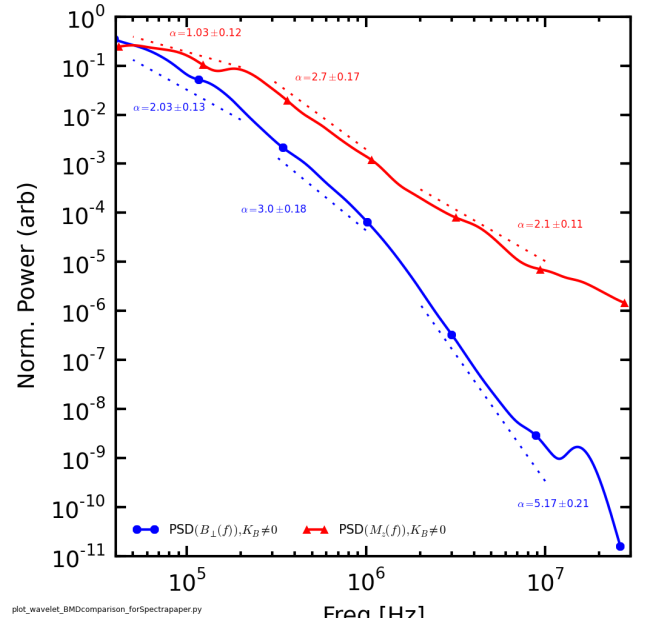


FIG. 6. Comparison of magnetic power spectra (blue circles) and Mach number power spectra (red triangles) for the non-zero helicity state as a function of frequency. The curves are normalized to their peak values to highlight the comparison of their shapes so relative fluctuation power between the two is not indicated. Fits are made and their values and ranges are indicated in Table III.

sibly due to a dissipation mechanism that may be further tapping magnetic energy, and given the increase in velocity spectra scaling, even further adding to velocity fluctuation energy. It may, however, also be due to the frequency of response of the Mach probe itself. Alternative velocity fluctuation measurements are being considered (i.e. electric field fluctuations) for future runs for comparison.

A comparison of density fluctuation spectra from the line-integrated interferometer timeseries is also shown in Figure 5(c). Note that the fluctuation power of the $K_B = 0$ state is higher in this case. The spectra curves of the density tend to be the least power-law like, so deriving any conclusions from the data at this point are difficult, though there is slight evidence for a brief rise in the spectra corresponding to ion inertial length scales (about 1-3MHz). This is discussed further in Section VIII. Future work will make attempts to get a more localized measurement as well as seek correlations between magnetic and density fluctuations which can have implications for the underlying nature of the turbulence⁴¹.

VI. WAVENUMBER SPECTRA

A unique turbulence measurement that can be made in the SSX plasma is a direct wavenumber spectrum using a multi-tipped magnetic probe that is inserted radially into the wind-tunnel. The probe can measure $\vec{B}(t)$ at 16 locations along a 6.8cm length of the radius at a spacing

of 0.46cm. In Fourier space, this allows measurements of scales from about 7cm to 1cm. Given that the injection scale of the magnetic energy is on the order of the initial size of the spheromaks—15.5cm—and a dissipation scale can be estimated to be just under 1cm—ion inertial length for a $1.5 \times 10^{15} \text{cm}^{-3}$ plasma is on the order of 0.6cm—the spatial range sampled by the probe is most likely in the inertial range of the turbulent cascade. Note that k and $1/\lambda$ (wavenumber and inverse length) is used interchangeably here, with both intended to mean inverse length scale (i.e. 2π dropped).

Since the probe can take simultaneous measurements of \vec{B} across the plasma, snapshots of the spatial structure of the plasma can be made at each time step. In turn, these spatial distributions can be Fourier transformed to produced power-spectra of the scales. Thus, this measurement can capture the direct wavenumber spectra of the plasma turbulence without reliance on any Doppler shifting as is needed to invoke the Taylor Hypothesis. Moreover, since \vec{B} is constructed from three orthogonal measurements, the power-spectra of vectors perpendicular and parallel to the axial flow of the plasma can be separately analyzed.

The downside, of course, to a measurement of the wavenumber spectra in this way is the lack of resolution compared to a Doppler-shifted frequency spectrum. With only 16 spatial points, the Fourier spectrum can have only 8 points, and only seven can be displayed in log-log format. However, despite this lack of resolution, it is a more physically direct measurement and can be used to cross-reference other observations of spectra.

Figure 7 shows the wavenumber power spectrum for the two helicity states and for full-vector(a), perpendicular vector(b) and parallel vector(c), averaged for each timestep in between 40 and $60\mu\text{s}$ and over 40 shots. Comparison of the curves in Figure 7(a) seem to show a slight variation in the slopes between helicity states, though with the large given errors in the fit (due to low resolution), the slopes of the two curves are essentially the same. A similar trend is observed amongst the three curves in (b) and (c) as well. A larger difference arises when comparing the curves of different vectors. Namely, the separately computed perpendicular and parallel spectra in (b) and (c) tend to be slightly steeper than the full vector spectra in (a). Moreover, the parallel curves appear to be steeper than the perpendicular curves. Again, the error in the fits are large for the low resolution data, but the trends are very suggestive of a wavenumber anisotropy (rather than variance anisotropy). However, there is not as much conclusive evidence for this form of anisotropy as there is for the variance anisotropy.

The wavenumber spectra and frequency spectra can be directly compared by invoking the Taylor hypothesis for the frequency spectra and Doppler-shifting the frequency spectra by the bulk plasma velocity,

$$B(f) \longrightarrow B(f - kV_p) \longrightarrow V_p B(k) \quad (4)$$

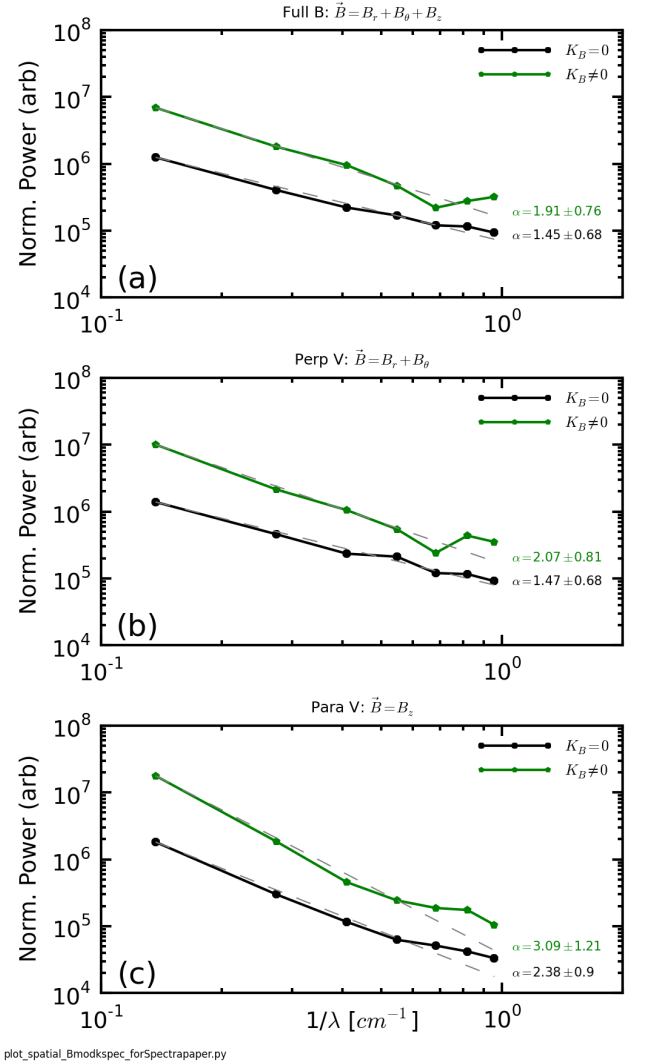


FIG. 7. Magnetic wavenumber power spectra constructed using a multi-point radially oriented probe for both $K_B \neq 0$ (green pentagons) and $K_B = 0$ (black circles) helicity states. The three panels show this spectrum for (a) the magnitude of the full magnetic field vector, $\vec{B} = (B_r, B_\theta, B_z)$, (b) the magnitude of the projected vector perpendicular to the axial flow, (B_r, B_θ) , and (c) the magnitude of the vector parallel to the axial flow, B_z . Power-law fits and errors are indicated at the right side of each curve and are fit over the entire data range shown.

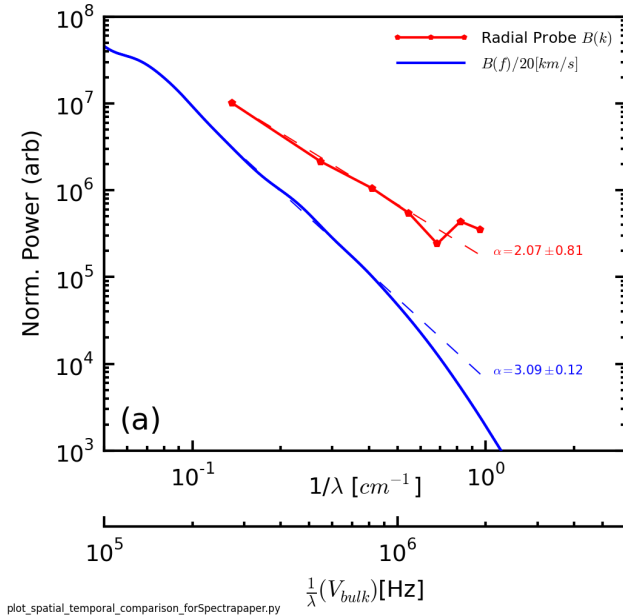


FIG. 8. Comparison of direct wavenumber spectra (red, pentagons) and frequency spectra (blue) shifted into spatial scales using the Taylor Hypothesis and a bulk flow of 20km/s as a function of inverse length scale. Fits are made to the same range in inverse length space and are indicated to the right of each curve. The double x-axes show the relative correspondence between inverse length scales and frequencies given the Taylor Hypothesis shift.

Typically, the last simplification in Equation 4 can only be done if the bulk flow is high enough such that $kV_p \gg f$ where f is the temporal frequency of any modes in the plasma. However, in this plasma, the nature of the present modes is still under investigation so for a first approximation comparison, the full Taylor Hypothesis transformation is used. For this plasma, the bulk velocity can be estimated (using both time-of-flight and Mach probe measurements) to be about 20km/s. The frequency spectra can be plotted on the same scale as the wavenumber spectra. Figure 8 shows this comparison for the non-zero helicity plasma. The bottom axes show under Taylor Hypothesis, the range in scales (10cm to 1cm) corresponds to frequencies of 200kHz-2MHz. The curves are placed arbitrarily on the y-axis. Maximum likelihood estimation power-law fits are made to the same range in both curves. The slopes of the curves are comparable suggesting that invoking the Taylor hypothesis for the frequency spectra is not entirely unwarranted. Instead, the steeper slope of the frequency spectra could be reflective of the effect of a combined temporal *and* spatial scaling, which the direct wavenumber spectrum does not include. However, breakdown of Taylor Hypothesis has been shown to make the spectra more shallow than steeper (Klein Taylor Hypothesis draft). The differences might also be reflective of a wavenumber anisotropy as the direct wavenumber spectra probes k_r and the Doppler-shifted frequency spectra probes k_z , similar to what is hinted in figure 7. Future experiments will seek

TABLE IV. Indices from MLE power-law fits of magnetic fluctuation spectra from experiment and simulation in Figure 9.

Fit Range [MHz]	Parameter	Index	Error
0.1-1.0	Sim: B_{\parallel}	3.22	0.11
	Sim: B_{\perp}	3.26	0.11
	Exp: B_{\parallel}	2.91	0.10
	Exp: B_{\perp}	2.72	0.10

to make a more direct comparison by aligning a multi-tipped probe in line with, rather than perpendicular to, the flow, in order to further probe the validity of the Taylor Hypothesis in this case.

VII. COMPARISON WITH SIMULATION

Simulations of the plasma produced in the SSX wind-tunnel have been generated with the HiFi spectral-element multi-fluid modeling framework using a set of normalized compressible Hall-MHD equations. Favorable comparisons of turbulent spectra and intermittency have been reported between these simulations and the experimental plasma²⁵. Further analysis is presented here which show similar observations of anisotropy, wavenumber spectra, and velocity-B-field spectra comparisons as is seen in the experiment. The simulation presented here has parameters most closely resembling the $K_B \neq 0$ helicity state. Details of the simulation equations and framework have been previously reported²⁵.

Timeseries of quantities in 3mm spheres approximately 1cm off the central axis and at the midplane are extracted from the simulation for density, three axes of magnetic field, three axes of velocity. To provide some ensemble averaging, points at eight different azimuthal angles are used since only one iteration of the simulation discharge has been produced.

The simulation timeseries data is analyzed in a similar manner as the experimental data with the exception that the mother wavelet used for the wavelet transform of the simulation data is a fourth-order Paul rather than a sixth-order Morlet, in order to better capture time resolution for the lesser sampled simulation. A variance anisotropy analysis is conducted in the same manner as well using a local magnetic field and the projection method. Figure 9(a) shows simulation decomposition into perpendicular and parallel spectra compared to $K_B \neq 0$ experimental spectra. The simulation and experimental spectra are staggered on the y-axis to emphasize features of the shape. Clearly, the simulation data exhibits growing variance anisotropy with increasing frequency. The slopes of the simulation spectra match qualitatively well in the region of 100kHz to 1MHz, though the fit spectra indices indicate a slightly steeper slope than the experiment (see Table IV). The high frequency end of the simulation spectra drops in power faster than the experiment, likely due to the limits in time resolution.

The trend in anisotropy ratio is also similar in the simulation and the experiment, as seen in Figure 9(b),

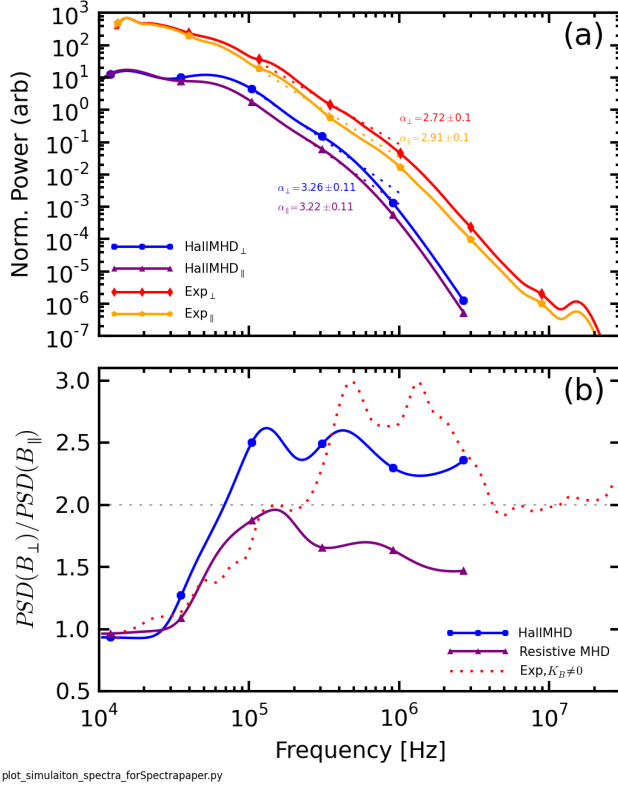


FIG. 9. (a) Magnetic power spectra for simulation (blue circles, purple triangles) and the non-zero helicity experiment (red diamonds, orange pentagons) as a function of frequency. The experimental and simulation curves are arbitrarily placed on the y-axis, but the perpendicular (upper) and parallel (lower) curves within each pair are placed relative to each other. Fits and fit regions are indicated in Table IV. (b) Anisotropy ratio as a function of frequency for a Hall-MHD simulation (blue circles), a resistive-MHD simulation (purple triangles) and the non-zero helicity experimental data (dotted red). An isotropic ratio is indicated by the horizontal dotted gray line.

though the simulation does not reach as large a peak ratio, nor does it scale at quite the same rate. Nevertheless, the clear observation of increase in ratio suggests that the compressible Hall-MHD physics captures the generation of the anisotropy. Figure 9(b) also shows the anisotropy ratio for a simulation run with the Hall term in the compressible MHD equations set to zero (i.e. resistive MHD). Unlike the Hall MHD and the experiment, the ratio does not switch over to perpendicular dominance, and instead stays near or below the isotropy line. The implications of this have not been analyzed in depth.

A comparison between velocity and magnetic field fluctuations in the simulation can also be made. Figure 10 shows wavelet transformed frequency power spectra for the total magnetic field (sum of B_r , B_{θ} , and B_z), total velocity, and density, all normalized to their respective peaks. Qualitatively, the comparison between velocity and magnetic field spectra supports the results of the experimental data for non-zero helicity plasmas: the peak

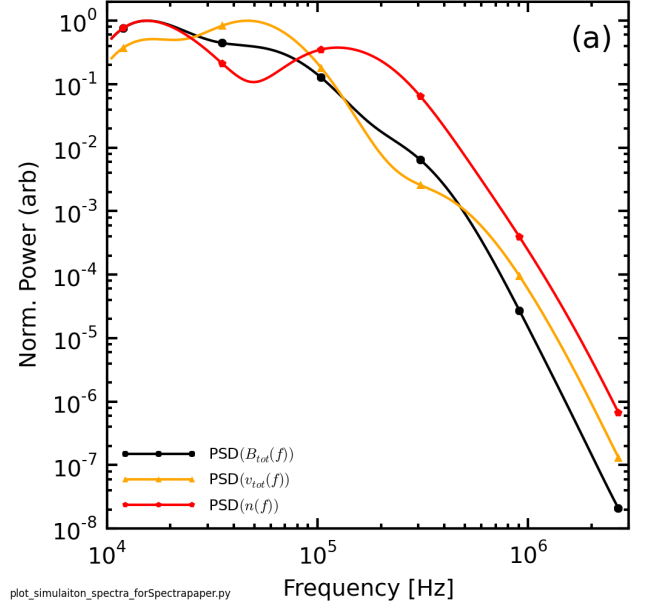


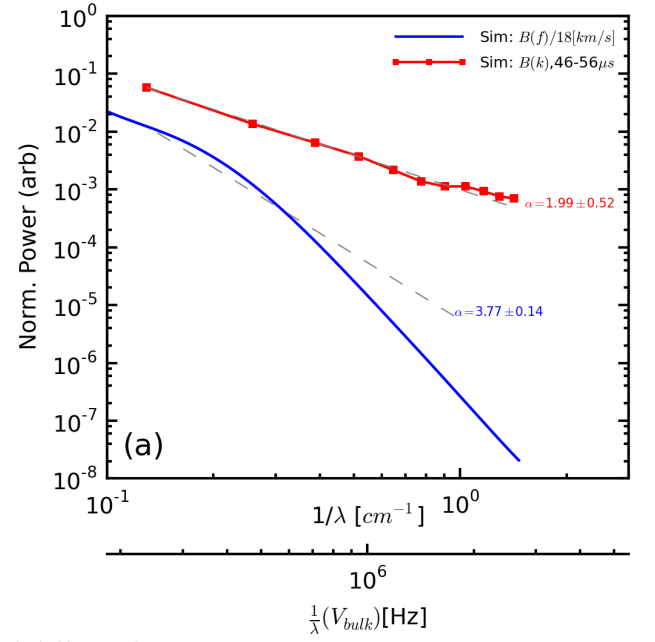
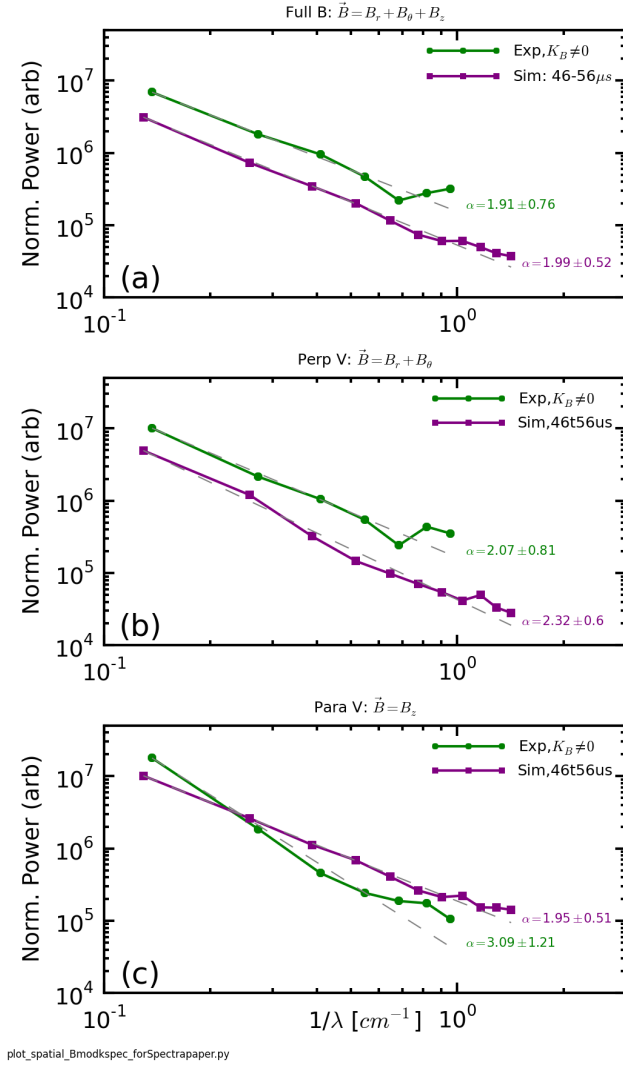
FIG. 10. Comparison of simulation generated magnetic (black circles), velocity (orange triangles) and density (red pentagons) power spectra as a function of frequency.

in the velocity spectra occurs at a larger frequency than the magnetic spectra. This suggests that energy for the velocity fluctuations are being injected at a smaller scale than the magnetic field fluctuations. Though not conducted here, further analysis of the simulation could potentially show direct energy transfer between the magnetic and velocity fluctuations.

The wavenumber spectra from a radial cut of the simulation data is also generated using 24 points. Figure 11 shows a direct comparison between simulation and experimental wavenumber spectra. The slightly higher spatial resolution allows the simulation to reach a smaller scale than the experiment, to about 0.7cm. In general, the comparison between simulation and experiment is good suggesting that the simulation is capturing well the spatial structure of the turbulence. Even though the simulation can observe slightly smaller scales, it does not appear to probe small enough to exhibit any dissipation effects at to ion inertial length scales which for the simulation is at about 0.7cm.

Like in the experiment, the spatial and temporal spectra of the simulation is compared and shown in Figure 12. The simulation has a bulk axial flow of 18km/s, close to the 20km/s observed in the experiment. A similar trend is seen with the spatial spectra having a shallower slope than the Doppler-shifted frequency spectra. The main difference again appears to be that the frequency spectra hits the limits of the temporal resolution at lower frequencies than the experiment.

These further comparisons of turbulent statistics and characteristics between the experimental plasma and a compressible Hall-MHD simulation help validate the model as useful for understanding the physical processes. Subsequent simulation analysis will likely entail more de-



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FIG. 12. Similar comparison of wavenumber and converted frequency spectra for simulation data as made in Figure 8, though for the simulation the bulk flow is 18km/s so the correspondence between the two x-axis scales is slight different than the previous plot.

tailed computation of how energy might be being distributed and moved through the plasma including relationships between magnetic field and velocity as well as between magnetic fluctuations perpendicular and parallel to a local B-field. Effects of varying helicity can also be explored.

VIII. DISCUSSION AND EVIDENCE FOR ION SCALE EFFECTS

A major remaining question for this analysis is whether the plasma diagnostics are able to observe effects of a dissipation scale in this turbulence. Perhaps, a more general question can be posed in how well does this plasma exhibit a traditional fluid-turbulence-like picture which posits an injection scale, inertial scale and dissipation scale.

The results of this analysis provide a number of hints that the ion inertial scale length is being probed, but no one piece of evidence is strong enough to make a conclusive assertion. The first clue arises by comparing the break-point of the magnetic spectra with the Doppler-shifted ion inertial scale length, $f_{\delta i}$. Figure 13(a) shows the spectra for magnetic field, Mach number and density for the $K_B \neq 0$ data all normalized to their respective maximum value, with dashed lines indicating the Doppler-shifted frequency of ion inertial length, $f_{\delta i}$ using a bulk flow speed of 20km/s, the collision frequency, ν_i , and ion cyclotron frequency, f_{ci} . Note that

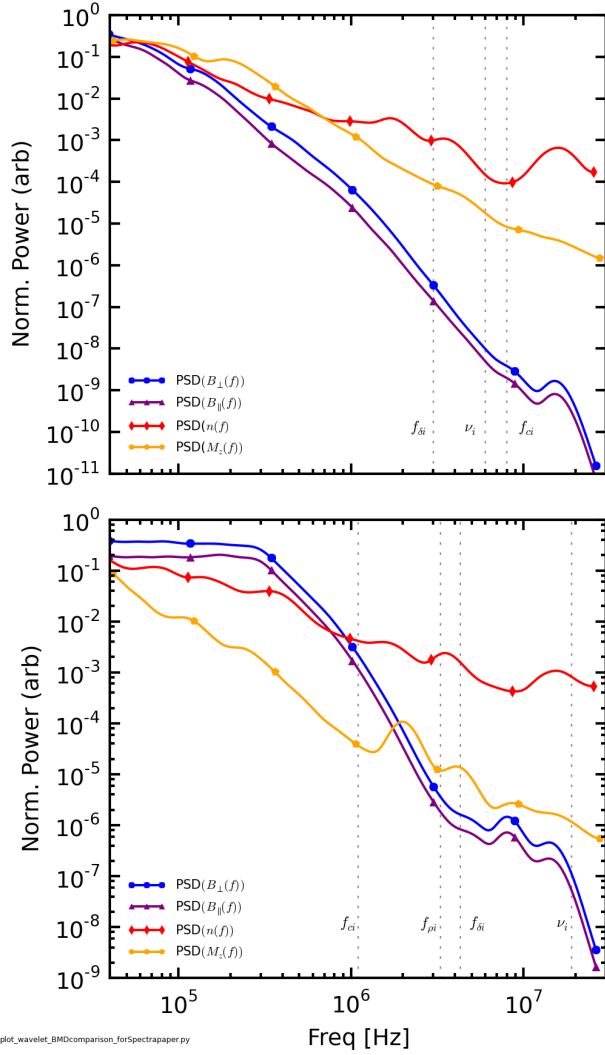


FIG. 13. (a) Comparison of magnetic(perp-blue circles,para-purple triangles), Mach(orange pentagons), and density(red diamonds) spectra all normalized to their respective maximum value to highly differences in spectral shape for $K_B \neq 0$ data. Vertical dashed lines indicate the position in frequency space of the Doppler-shifted ion inertial length, $f_{\delta i}$, the collision frequency, ν_i , and the ion cyclotron frequency, f_{ci} . (b) Similar comparison as (a), but for the zero helicity case. The Doppler-shifted ion gyroradius, $f_{\rho i}$ is also indicated on this second subplot.

the break point occurs just before the ion inertial frequency is reached. Since the ion inertial scale is often associated with the scale size of reconnection layers or current sheets, a break-point just proceeding this scale suggests the onset of a dissipation mechanism associated with current sheets of some form(citation needed).

Two supporting pieces of evidence for this hypothesis come from the comparison of the density, flow and magnetic spectra as well as the trend in variance anisotropy. The red curve in Figure 13(a) shows a slight bump up again around $f_{\delta i}$. This is possibly evidence of the density bump effect observed when the plasma becomes more compressible—which would be expected at ion inertial length scales. The flattening of the velocity spectrum has also been speculated to be a signal of compressive effects³⁶, but no clear cut relationship has as yet been observed. Similarly, the anisotropy ratio in Figure 2(b) begins to decrease at about this frequency (see Figure 3) which also suggests an increase in compressibility. Such decreases in anisotropy ratio at dissipation scales have been observed in the solar wind^{27,37}.

However, there remain other explanations for these effects. One, the Taylor Hypothesis used to establish the connection between the frequency and the scale length here is not obviously valid. If the Taylor Hypothesis cannot be invoked, there may be other reasons for the break-point observed. Also, the shallowing of the spectra in Figure 13(a) may be due to reaching the noise floor set by the bit-depth of the interferometer/Mach probe data acquisition system (smaller than the 14-bit range of the magnetic). Moreover, there is evidence, at least in the solar wind, that density spectra might be expected to steepen beyond ion scales⁴², not flatten.

Comparison of the two different helicity states only adds to the ambiguity. Figure 13(b) shows the same curves as in (a), but for the zero-helicity state. The break-point in the magnetic field here appears to occur close to the ion cyclotron frequency rather than the ion inertial length. It should be noted though that previous work²⁶ has suggested that the zero-helicity state consists of much fewer current sheets and as such, dissipation in this state would depend much less on those mechanisms. The zero-helicity state also shows no change in anisotropy with scale nor any density bump.

The results presented, nevertheless, do highlight the need for further investigation, particularly in whether a Taylor Hypothesis can be invoked, or the very least, whether some type of Doppler shift can be applied to properly connect the frequency and scale size of the signal. Similarly, a higher resolution in the spatial probes could provide confirmation as its current state just misses the apparent dissipation scale. Moreover, better resolved density and Mach flow diagnostics would be useful in order to distinguish between noise and real effect.

Lastly, the simulation could potentially provide some insight into the processes occurring. Unfortunately, the simulation fails to probe at a small enough scale to show ion scale effects. Both the frequency spectra and

wavenumber spectra fall short ion inertial scales. The variance anisotropy ratio supports this as an increase in the ratio is observed, but not a decrease at smaller scales/higher frequencies. However, the favorable comparison between the experiment and the simulation at smaller frequencies and larger scales means that improvement of the simulations resolution maybe be useful for exploring the dissipation physics.

IX. CONCLUSIONS

This manuscript presents a broad analysis of both temporal and spatial fluctuation spectra with the intent of understanding the MHD turbulence observed in a wind-tunnel on the Swarthmore Spheromak Experiment. The results show the presence of variance anisotropy—more power in fluctuations perpendicular to the local magnetic field than parallel—for a plasma with non-zero injected helicity. The ratio of perpendicular to parallel shows variation as a function of frequency and of time, reaching peak values of about $R=3$ during the period of mostly stationary fluctuations and at frequencies between 500kHz and 1MHz. Very little anisotropy is observed for a zero-helicity state, which maybe be more closely related to the value of β for this state of the plasma rather than the helicity itself.

Comparison of fluctuation spectra of magnetic fields and velocity appear to support the notion that energy is primarily injected into the plasma through magnetic fields (during the spheromak formation process). Comparison amongst magnetic, velocity, and density spectra provide clues for the nature of the dissipation processes occurring in the plasma, but much more research is needed for conclusions to be made.

A direct wavenumber spectrum is also measured in the plasma, highlighting an advantage of measurements in laboratory plasmas over *in-situ* space measurements. The wavenumber spectra show slight differences when compared to temporal frequency spectra converted into spatial scales using the Taylor Hypothesis and the measured bulk axial flow of the plasma. Wavenumber spectra also suggest a wavenumber anisotropy, but the results are less conclusive than the observation of the variance anisotropy, but the results do set-up for future analysis of the validity of the Taylor Hypothesis.

All of these results are also compared to a Hall-MHD computation generated to simulate the spheromak relaxation process. Turbulent spectra using synthetic diagnostics compare favorably to the experiment particularly in the measured spectral indices for power-law fits. The simulation does not probe as small a frequency or spatial scale as the experiment, however.

Many open questions still remain though, including whether or not dissipation at the ion inertial length scale is observed. Both the magnetic spectra breakpoint and a decrease in anisotropy occur at scales associated with the ion inertial length, but only when the Taylor Hypothesis

can be invoked. The direct wavenumber spectra do not probe a smaller enough scale to provide any evidence for or against this hypothesis. The encouraging comparisons with simulation, though, might provide some support if it can be pushed to higher resolutions. Other diagnostic techniques are also being pursued to help unravel this issue.

The goal of this analysis is to establish the use of the MHD wind-tunnel as a testbed for understanding turbulence typically researched at space physics scales. Many of these results have intriguing comparisons with situations in the solar wind turbulence for example: both cases where the data compare well, and where they do not. The advantage of laboratory experiments lies in the ability to make detailed spatial measurements, the ability to have some nominal control over parameters, and the ability to make many repeatable measurements. Through a combination of these laboratory observations, computer simulation *and* space observation, a greater understanding of MHD turbulence is sought.

Appendix A: Computation of Variance Anisotropy

The computation of the variance anisotropy involves two main steps: a determination of fluctuation power and an estimate of the distribution of the fluctuation power relative to a vector direction. The first step is accomplished using a Wavelet transform procedure as is discussed in the text. The division of fluctuation power is accomplished using what will be described here as a projection method. A second estimate of fluctuation power distribution is constructed using a threshold method and is also described in detail in this appendix. The threshold method is a more straightforward process for determining the level of variance anisotropy, but suffers from diminishing resolution. Its presentation here is mainly as a validation of the eventual use of the projection method for the analysis in the paper.

1. Threshold Method

Both the threshold method and the projection method for determining variance anisotropy rely on the ability of the wavelet transform to yield a power spectrum distribution (PSD) as a function of both time and frequency (i.e. $B(f,t)$). A local (in time) vector \vec{B} is determined for each time, t ,

$$\vec{B}(t) = B_r(t)\hat{r} + B_\theta(t)\hat{\theta} + B_z(t)\hat{z} \quad (A1)$$

where each component $n = r, \theta, z$ is determined from \dot{B}_n by integrating over time as

$$B_n(t) = \int_0^t \frac{d}{d\tau} B_n(\tau) d\tau. \quad (A2)$$

Since the magnetic probe measures orthogonal magnetic field directions by construction, this fact can be

used to directly seek a difference in fluctuation spectra depending on orientation perpendicular or parallel to the overall field. Thus, a threshold ratio for each component can be defined as a fraction of the total magnitude as in,

$$R_n \geq \frac{|B_n|^2}{|\vec{B}|^2} \quad (\text{A3})$$

which reflects the relative amount that the total vector points in one of the three orthogonal directions. Then for every time, t , in a given time range and for each shot, the quantity $B_n(f, t)$ is summed for each t where

$$\frac{|B_n|^2}{|\vec{B}|^2} \geq R_n \quad (\text{A4})$$

for all frequencies, f . The value chosen for R_n determines how strictly total vector aligns with the component, n . By definition, then, the summed power for B_n is considered the parallel component and the sum of the remaining two directions is the perpendicular component. If there is any anisotropy in the signal, a difference in energy content of the spectra should become apparent as the threshold value is increased.

Indeed, an effect like this is observed. Figure 14 shows the ratio of total perpendicular fluctuation power to parallel fluctuation power for $n = r$ (a) and $n = \theta$ (b). The threshold fraction is indicated by color. The dashed line at 2 represents isotropy—where the sum of 2 perpendicular components is about twice the power of the single parallel component. Clearly, for the lowest threshold value, the ratio remains close to the isotropy line for all frequencies as would be expected. As the threshold value is raised, the ratio from about 10kHz and higher begins to grow. This shows there is variance anisotropy in the plasma. If the plasma were isotropic, a difference between perpendicular and parallel spectra would not be seen. The anisotropy ratio reaches a maximum as the threshold nears 100%. The drawback to this method, however, is that as the threshold is increased, the number of individual spectra summed is reduced which increases the error of each curve. This also probably explains why the threshold method indicates higher ratios than the more highly averaged data from the projection method discussed next.

2. Projection Method

An alternative method, and that used to compute the results presented in this manuscript, uses the \vec{B} time-series data to project spectral power into perpendicular and parallel portions at each timestep. This projection method uses all the available timesteps and shots rather than making a cut like the threshold method. It will be shown later that the two methods give quantitatively similar answers for the amount of variance anisotropy.

The projection method also uses the Wavelet transform to compute $B(f, t)$. However, that than use the $\vec{B}(t)$ as a

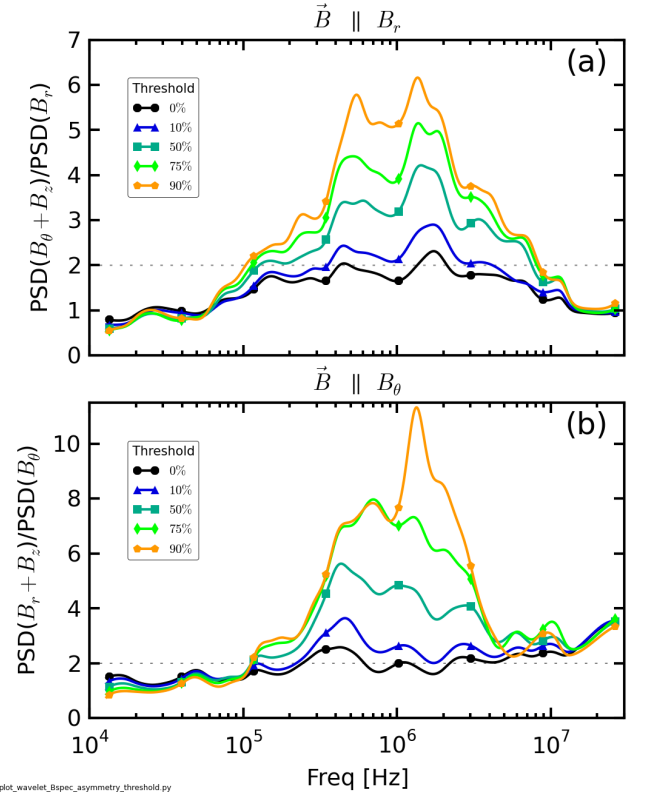


FIG. 14. (a) Ratio of the sum of magnetic power spectra of B_θ and B_z over the spectra of B_r as a function of frequency and summed over the number of instances in time when the threshold ratio, defined in Equation A3 for $n = r$, is surpassed. Each curve indicates the level of anisotropy for different threshold values. The threshold ratio indicates the relative amount the full magnetic field vector, \vec{B} points in the direction of B_r . (b) Similar to (a), but for $n = \theta$.

threshold value, it is used as a reference vector to determine what fraction of the fluctuation power of $B_r(f, t)$, $B_\theta(f, t)$, and $B_z(f, t)$ is perpendicular or parallel to that vector. The parallel component of each $B_x(f, t)$ is found by computing the projection,

$$Proj_u v = \frac{\vec{v} \cdot \vec{u}}{|\vec{u}|} \vec{u} \quad (\text{A5})$$

which shows that the magnitude of the component of B_x parallel to \vec{B} is

$$(B_x^\parallel)^2 = \frac{B_x^2}{|\vec{B}|^2}. \quad (\text{A6})$$

Then, the magnitude of the component perpendicular is

$$(B_x^\perp)^2 = 1 - \frac{B_x^2}{|\vec{B}|^2}. \quad (\text{A7})$$

Using these projection coefficients, each Wavelet transform spectra, $B_x(f, t)$ for $x = r, \theta, z$, can be divided into $B^\parallel(f, t)$ and $B^\perp(f, t)$. For each timestep during each shot, the total parallel and perpendicular power is found

by summing the respective portions from each orthogonal direction, as in

$$B_{\parallel}(f) = \sum_{t=t_0}^{t_1} (B_r^{\parallel})^2(B_r(f, t)) + (B_{\theta}^{\parallel})^2(B_{\theta}(f, t)) + (B_z^{\parallel})^2(B_z(f, t)) \quad (\text{A8})$$

$$B_{\perp}(f) = \sum_{t=t_0}^{t_1} (B_r^{\perp})^2(B_r(f, t)) + (B_{\theta}^{\perp})^2(B_{\theta}(f, t)) + (B_z^{\perp})^2(B_z(f, t)) \quad (\text{A9})$$

for the given time range $t_0 \leq t \leq t_n$. These summed quantities, $B^{\parallel}(f, t)$ and $B^{\perp}(f, t)$ are what are used in Section IV.

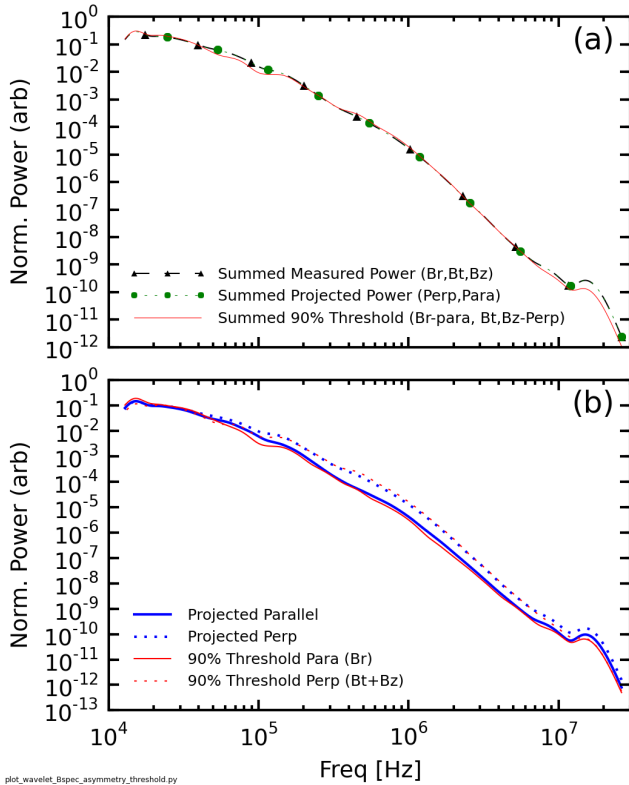


FIG. 15. (a) Comparison of the total summed magnetic power showing the exact correspondance between (B_r, B_{θ}, B_z) and $(B_{\parallel}, B_{\perp})$. Red curve shows the total power for the 90% threshold level which while not exact, is reasonably close to the other two curves. (b) Comparison of anisotropy in magnetic spectra using the projected method and the threshold method.

As a check, the total fluctuation power spectrum is computed in three different ways and plotted in Figure 15(a) for a time range of 40 to $60\mu\text{s}$. The total power is found by (1) summing $B_r(f)$, $B_{\theta}(f)$, and B_z directly, (2) summing $B^{\parallel}(f)$ and $B^{\perp}(f)$, (3) and summing the 90% threshold spectra of B_r , B_{θ} , and B_z from the

threshold method. The curves are averaged over the total number of timesteps used in their construction so they can be directly compared amongst one another. The first two ways match exactly, showing that the total power is being properly portioned. The 90% threshold calculation is does not match exactly, though is close. Figure 15(b) shows a comparison of the variance anisotropy of the frequency power spectra as computed by the threshold and the projection methods. Again, the curves are normalized to total number of timesteps used in construction. The quantitative comparison shows that the projection method works well to compute the level of anisotropy as it compares well to the more direct, robust threshold method.

Appendix B: Wavelet vs. FFT

Figure 16 demonstrates the correspondance between magnetic fluctuation spectra using the Wavelet analysis discussed in Section III and a traditional Fast Fourier Transform. The red curves in both Fig 16(a) and (b) are the Wavelet generated spectra for the zero and non-zero helicity states respectively. The gray curves show a Fourier transform generated spectra for only the $40\text{-}60\mu\text{s}$ range of each shot. Clearly, the overall shape between the two sets of curves is nearly identical. Deviations occur at low frequencies where the Wavelet transform can sample slightly lower frequencies since the analysis uses the entire time range rather than a time subset. The nature of the Wavelet transform also allows for higher resolution binning especially at these lower frequencies. The Wavelet transform, especially with the particular mother wavelet used—Morlet—tends to cause some smoothing in frequency space compared with the FFT. This is clearly seen at higher frequencies as a modes around 10MHz are more clearly observed in the FFT curver than the wavelet. However, for this particular dataset, these modes are not pertinent as they are caused by characteristics of the gun system and not by turbulence physics.

Appendix C: Local vs. Global Magnetic Field

The use of a temporally local versus global magnetic field in anisotropy analysis in the solar wind has been often debated^{38,39}. In this paper, a local magnetic field reference vector has been used, but a relative global field can also be used to establish anisotropy. Since the experimental data is extracted on a shot-by-shot basis, the global field in this case is the mean field for the time duration of each shot. Figure 17 shows the anisotropy ratio for local field (reference vector at each timestep) and global field (mean field for each shot). While the local field yields a ratio that is slightly higher than for the global, the trend as a function of frequency is clearly similar. This comparison can be extended to the FFT anal-

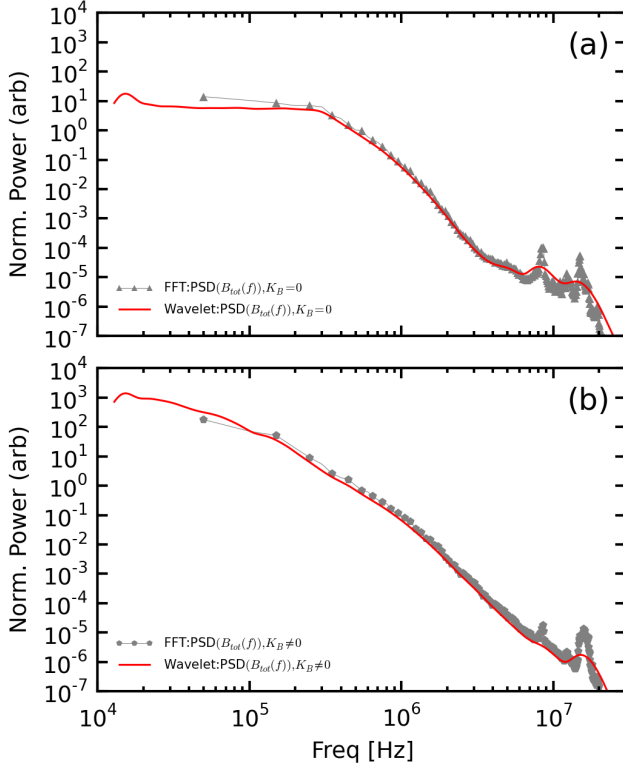


FIG. 16. Comparison of magnetic power spectra constructed using a Wavelet transform (red) and an FFT (gray triangles) for non-zero helicity (a) and zero helicity (b) states.

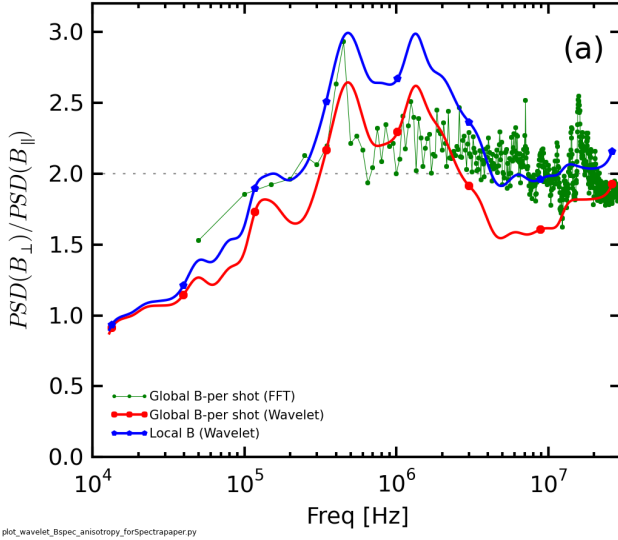


FIG. 17. Comparison of anisotropy ratio for magnetic power spectra constructed when using either a local (blue pentagon) or global (red circle) definition of the mean field. The thin green curve shows the anisotropy with a global mean field using an FFT transform rather than Wavelet.

ysis which typically would not be useful for a variance analysis technique because it does not have the time resolution that the Wavelet transform does. However, since the magnetic field does not change too quickly in these plasmas, a mean field can be chosen for each shot and then used to project the full FFT spectra generated for each shot. The green curve shows this ratio, which while not as distinct as the wavelet generated ratio, nevertheless shows a similar trend especially in the kHz to MHz range. These results all suggest that though the use of a global versus local field may modify the exact numerical relationship, an anisotropy trend can be observed in either case if it is present in the plasma.

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