

Morphology
of LMC and
SMC

Daniel
Wysocki

Magellanic
Clouds

Variable
Stars

Galactic
Morphology

Morphology of the Large and Small Magellanic Clouds using Fundamental Mode Cepheids

Rochester Academy of Science 2014

Daniel Wysocki

November 15, 2014

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2 Variable Stars

3 Galactic Morphology

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Magellanic Clouds

What are they?

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- two dwarf galaxies which orbit the Milky Way

What are they?

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- two dwarf galaxies which orbit the Milky Way
- irregular galaxies

Large Magellanic Cloud (LMC)

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- 50kpc away

Small Magellanic Cloud (SMC)

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- 60kpc away

Why are they important?

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- gravitationally interacting with our own galaxy

Why are they important?

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- gravitationally interacting with our own galaxy
- nearby galaxies

Why are they important?

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- gravitationally interacting with our own galaxy
- nearby galaxies
 - can be observed in detail

Why are they important?

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- gravitationally interacting with our own galaxy
- nearby galaxies
 - can be observed in detail
 - can be used as a distance calibrator to more distant galaxies

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Variable Stars

What are they?

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- stars whose luminosity changes with time

What are they?

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- stars whose luminosity changes with time
- many different types

What are they?

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- stars whose luminosity changes with time
- many different types
 - Classical Cepheids, Type II Cepheids, RR Lyrae, MIRA Variables, Delta Scutis, and more

What are they?

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- stars whose luminosity changes with time
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- some vary periodically

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 - Classical Cepheids, RR Lyrae, and Delta Scutis to name a few

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- stars whose luminosity changes with time
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 - Classical Cepheids, Type II Cepheids, RR Lyrae, MIRA Variables, Delta Scutis, and more
- some vary periodically
 - Classical Cepheids, RR Lyrae, and Delta Scutis to name a few
- the variation can be related to physical properties of the star

Light curves

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- function of a star's brightness over time

Light curves

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- function of a star's brightness over time
- for periodic variables, time can be transformed into phase

Light curves

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- function of a star's brightness over time
- for periodic variables, time can be transformed into phase
- shape, amplitude, and period of a star's light curve can reveal many things

Classical Cepheids

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- periodic variable stars

Classical Cepheids

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- periodic variable stars
 - period ranges from days to months

Classical Cepheids

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- periodic variable stars
 - period ranges from days to months
 - obey a period-luminosity-color relationship

Period-Luminosity-Color Relationship

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- luminosity depends on surface area and temperature

$$L \propto AT^4 \quad (1)$$

Period-Luminosity-Color Relationship

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- surface area and temperature cannot be measured directly

Period-Luminosity-Color Relationship

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- luminosity depends on surface area and temperature

$$L \propto AT^4 \quad (1)$$

- surface area and temperature cannot be measured directly
 - period of oscillation depends on size

Period-Luminosity-Color Relationship

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- luminosity depends on surface area and temperature

$$L \propto AT^4 \quad (1)$$

- surface area and temperature cannot be measured directly
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 - color depends on the temperature

$$\overline{M}_\lambda = \alpha_\lambda \log P + \beta_\lambda + \epsilon_\lambda \quad (2)$$

Period-Luminosity-Color Relationship

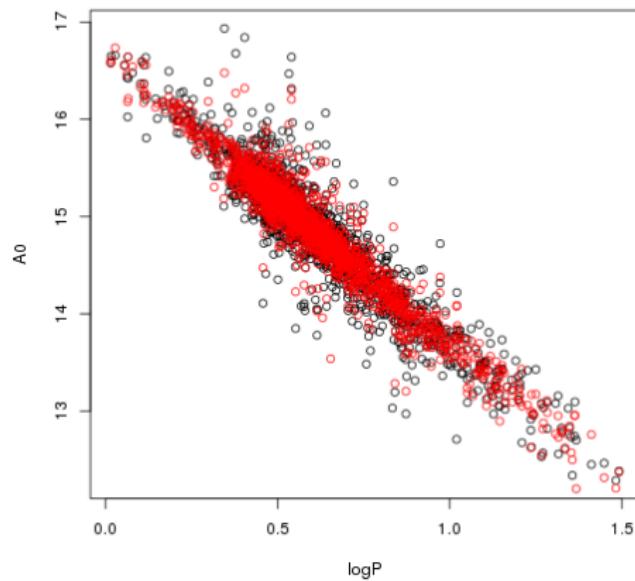
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Standard Candles

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- imagine a candle, whose luminosity is known

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- imagine a candle, whose luminosity is known
- place that candle at the other end of a field

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- imagine a candle, whose luminosity is known
- place that candle at the other end of a field
- by comparing the observed brightness to the known luminosity, the length of the field can be determined

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- imagine a candle, whose luminosity is known
- place that candle at the other end of a field
- by comparing the observed brightness to the known luminosity, the length of the field can be determined
 - intermediate gas and dust affects the observed brightness

Standard Candles

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- imagine a candle, whose luminosity is known
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 - intermediate gas and dust affects the observed brightness
 - this can be accounted for by modelling the gas and dust

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- imagine a candle, whose luminosity is known
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- by comparing the observed brightness to the known luminosity, the length of the field can be determined
 - intermediate gas and dust affects the observed brightness
 - this can be accounted for by modelling the gas and dust
- Classical Cepheids can be used in the same way

Distance Modulus

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- difference between apparent and observed magnitudes

$$\mu_i = m_i - M_i \quad (3)$$

Distance Modulus

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- difference between apparent and observed magnitudes

$$\mu_i = m_i - M_i \quad (3)$$

- substitute into equation from earlier

$$\overline{m}_{\lambda,i} = \alpha_{\lambda} \log P_i + \beta_{\lambda} + \mu_i + \epsilon_{\lambda,i} \quad (4)$$

Distance Modulus

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- difference between apparent and observed magnitudes

$$\mu_i = m_i - M_i \quad (3)$$

- substitute into equation from earlier

$$\overline{m}_{\lambda,i} = \alpha_{\lambda} \log P_i + \beta_{\lambda} + \mu_i + \epsilon_{\lambda,i} \quad (4)$$

- for N stars, this results in a system of $2N$ equations with $N + 4$ unknowns

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Galactic Morphology

Equatorial Coordinate System

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- right ascension (RA or α) is the astronomical equivalent of longitude

Equatorial Coordinate System

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- right ascension (RA or α) is the astronomical equivalent of longitude
- declination (Dec or δ) is the astronomical equivalent of latitude

Equatorial Coordinate System

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- right ascension (RA or α) is the astronomical equivalent of longitude
- declination (Dec or δ) is the astronomical equivalent of latitude
- these two angles can be used to describe an object's location in the sky

Equatorial Coordinate System

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- right ascension (RA or α) is the astronomical equivalent of longitude
- declination (Dec or δ) is the astronomical equivalent of latitude
- these two angles can be used to describe an object's location in the sky
- introduce distance (D), and the coordinate system now describes three dimensional space

Cartesian Coordinate System

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- familiar x, y, z coordinate system

Cartesian Coordinate System

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- familiar x, y, z coordinate system
- can be obtained from equatorial coordinate system through the following transformations

$$x = -D \sin(\alpha - \alpha_0) \cos \delta,$$

$$y = D \sin \delta \cos \delta_0 - D \sin \delta_0 \cos(\alpha - \alpha_0) \cos \delta,$$

$$z = D_0 - D \sin \delta \sin \delta_0 - D \cos \delta_0 \cos \alpha - \alpha_0 \cos \delta$$

LMC in 3D

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Inclination and Position Angles

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- describe the orientation of a galaxy with respect to Earth

Inclination and Position Angles

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- describe the orientation of a galaxy with respect to Earth
- can be obtained in different ways

Plane-fitting Method

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- most common method for obtaining inclination and position angles

Plane-fitting Method

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- most common method for obtaining inclination and position angles
- fits the 2D plane of best fit to the collection of stars

Plane-fitting Method

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- most common method for obtaining inclination and position angles
- fits the 2D plane of best fit to the collection of stars
- basic linear algebra is used to find the 2 angles

Plane-fitting Method

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- most common method for obtaining inclination and position angles
- fits the 2D plane of best fit to the collection of stars
- basic linear algebra is used to find the 2 angles
- does not do a very good job describing the 3D structure of the galaxy

Ellipsoid-fitting Method

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- less common method for obtaining inclination and position angles

Ellipsoid-fitting Method

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- less common method for obtaining inclination and position angles
- fits a 3D ellipsoid to the collection of stars

Ellipsoid-fitting Method

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- less common method for obtaining inclination and position angles
- fits a 3D ellipsoid to the collection of stars
- principal axis transformation is performed on moment of inertia tensor

Ellipsoid-fitting Method

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- less common method for obtaining inclination and position angles
- fits a 3D ellipsoid to the collection of stars
- principal axis transformation is performed on moment of inertia tensor
- transformation gives eigenvalues and eigenvectors, which are used to describe the size and orientation of the axes of the ellipsoid

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