α
α
α
α
?
S_ν
T_e
$S_\nu \propto$
$e^{-h\nu/kT_e}/\sqrt{T_e}$
$S_\nu =$
$2kT_e\nu^2/c^2$
κ_λ
$\kappa_\lambda \propto$
$\lambda^{2.1}$
τ_λ
$\tau_{20cm}/\tau_{0.7cm} \approx$
10^3
frozen-
in
α
α
??
im-
fit
im-
fit
??
??
$F_\nu 5\sigma$
im-
fit
$_1$999.*The flux density values used in the following Sections are the peak values listed in Table* **??**. *We assume absolute flux dens*
$_t$*emplate*/6/*atau*$_{qk}$*a.ps*[*trim* = 130*pt*50*pt*150*pt*10*pt*, *clip*, *width* = 7.5*cm*, *height* = 6.0*cm*]/*home*/*eamon*/*thesis*/*thesis*$_t$*er*
α
α
α
$\nu^a \lambda F_\nu$ *Imfit* *Imfit*
$^{-1}F_\nu$ F_ν $^{-1}$
α
α
^{a}*clean*
α
$>\sigma$
α
??
186
α
?
α
$F_K \propto$
ν^α
α
σ
$\pm$
μ
??
$\gtrsim\sigma$
3
α
5σ
|
|
α α
α
α
106
α
$>\sigma$
α
α
3σ
α
$_1$981*and Parkes slee*$_1$989 *radio telescopes, reported the detection of flares from single red giants. These transient radio event*
observed however, even with more sensitive interferometers, suggesting that such detections were spurious (e.g., $beasley_1$99
free emission from a luminosity class III single red giant at centimeter wavelengths was of α
$_1$983, $drake_1$986. *Since then there has been a modest number of centimeter and millimeter observations of this star. In Table* **??**
$\nu \leq$
250
$\nu \leq$
250
α ν F_ν
$\leq(3\sigma)$
α $\leq(3\sigma)$
$\leq(3\sigma)$
$\leq(3\sigma)$
??????
α
σ
$\sim\%$
σ
?
?
$_2$013. *Another possibility for the difference in values is that the longer cycle time used by* [**?**], *which was over double our value, m*
$_t$*emplate*/6/*aboo*$_s$*ummary.ps*[*Spectral energy distribution of* α
α
$\leq$
$\nu \leq$
?
$\sim$
$R_\star$
??
??
$_t$*emplate*/6/*atau*$_s$*ummary.ps*[*Spectral energy distribution of* α
α
$\leq$
$\nu \leq$
α
??