

α
 α
 S_ν
 T_ν^e
 $S_\nu^e \propto$
 $e^{-h\nu/kT_e}/\sqrt{T_e}$
 $S_\nu =$
 $2kT_e\nu^2/c^2$
 $\kappa_\lambda \propto$
 $\lambda^{2.1}$
 τ_λ
 $\tau_{20\text{cm}}/\tau_{0.7\text{cm}} \approx$
 10^3
frozen-
 α
 α
 $??$
 $im-$
 fit
 $im-$
 fit
 $??$
 $??$
 $F_\nu 5\sigma$
 $im-$
 fit
1999. The flux density values used in the following Sections are the peak values listed in Table ?? . We assume absolute flux densi-
 $t_{\text{emplate/6/atau}_qk.a.ps}$ [trim = 130pt 50pt 150pt 10pt, clip, width = 7.5cm, height = 6.0cm] /home/eamon/thesis/thesis.tex
 α
 α
 $\nu^a \lambda F_\nu \quad Imfit \quad Imfit$
 $^{-1} F_\nu \quad F_\nu \quad ^{-1}$
 α
 α_{clean}
 $>\sigma$
 α
 186
 α
 $F_\alpha \propto$
 ν
 $\alpha \pm$
 μ
 $??$
 $>\sigma$
 α
 σ
 α
 α
 α
1981 and Parkes 1989 radiotelescopes, reported the detection of flares from single red giants. These transient radio events
observed however, even with more sensitive interferometers, suggesting that such detections were spurious (e.g., Beasley 1999).
free emission from a luminosity class III single red giant at centimeter wavelengths was of α
1983, Drake 1986. Since then there has been a modest number of centimeter and millimeter observations of this star. In Table ?? ,
 $\nu <$
 250
 $\nu <$
 250
 $\nu \quad F_\nu$
 α
 $\leq (3\sigma)$
 α
 $\leq (3\sigma)$
 $\leq (3\sigma)$
 $\leq (3\sigma)$
 $?????$
 α
 σ
 $\sim \%$
 α
 α
2013. Another possibility for the difference in values is that the longer cycle time used by [?], which was over double our value, may
 $t_{\text{emplate/6/aboo}_s\text{ummary.ps}}$ [Spectral energy distribution of α
 α
 \leq
 $\nu <$
 α
 R_*
 $??$
 $??$
 $t_{\text{emplate/6/atau}_s\text{ummary.ps}}$ [Spectral energy distribution of α
 α
 \leq
 $\nu <$
 α
 $??$