

THE GALAXY UV LUMINOSITY FUNCTION BEFORE THE EPOCH OF REIONIZATION

CHARLOTTE A. MASON^{1,2}, MICHELE TRENTI³, AND TOMMASO TREU^{1,2}¹ Department of Physics, University of California, Santa Barbara, CA, 93106-9530, USA; cmason@physics.ucsb.edu² Department of Physics and Astronomy, UCLA, Los Angeles, CA, 90095-1547, USA³ School of Physics, University of Melbourne, Parkville, Victoria, Australia

Received 2015 August 5; accepted 2015 September 25; published 2015 October 22

ABSTRACT

We present a model for the evolution of the galaxy ultraviolet (UV) luminosity function (LF) across cosmic time where star formation is linked to the assembly of dark matter halos under the assumption of a mass-dependent, but redshift-independent, efficiency. We introduce a new self-consistent treatment of the halo star formation history, which allows us to make predictions at $z > 10$ (lookback time $\lesssim 500$ Myr), when growth is rapid. With a calibration at a single redshift to set the stellar-to-halo mass ratio, and no further degrees of freedom, our model captures the evolution of the UV LF over all available observations ($0 \lesssim z \lesssim 10$). The significant drop in luminosity density of currently detectable galaxies beyond $z \sim 8$ is explained by a shift of star formation toward less massive, fainter galaxies. Assuming that star formation proceeds down to atomic cooling halos, we derive a reionization optical depth $\tau = 0.056^{+0.007}_{-0.010}$, fully consistent with the latest Planck measurement, implying that the universe is fully reionized at $z = 7.84^{+0.65}_{-0.98}$. In addition, our model naturally produces smoothly rising star formation histories for galaxies with $L \lesssim L_*$ in agreement with observations and hydrodynamical simulations. Before the epoch of reionization at $z > 10$ we predict the LF to remain well-described by a Schechter function, but with an increasingly steep faint-end slope ($\alpha \sim -3.5$ at $z \sim 16$). Finally, we construct forecasts for surveys with *James Webb Space Telescope* (JWST) and *Wide-field Infrared Survey Telescope* (WFIRST) and predict that galaxies out to $z \sim 14$ will be observed. Galaxies at $z > 15$ will likely be accessible to JWST and WFIRST only through the assistance of strong lensing magnification.

Key words: cosmology: theory – galaxies: high-redshift – stars: formation

1. INTRODUCTION

The rest-frame ultraviolet (UV) galaxy luminosity function (LF) and its evolution with redshift are crucial tracers of galaxy properties over cosmic time. In particular, UV light can be used efficiently to measure the star formation rate (SFR), because photons at rest-frame wavelengths around 1500 Å are primarily produced by young, massive, and short-lived stars. Current observations characterize the UV LF over the large majority of the history of the universe, ranging from studies in the local universe from *Galaxy Evolution Explorer* data (e.g., Burgarella et al. 2006) to *Hubble Space Telescope Wide Field Camera 3* (HST/WFC3) observations that now reach redshift $z \sim 8$ –10, i.e., lookback times greater than 13 Gyr (e.g., Bouwens et al. 2015a). Transformational results on the rest-frame UV light from the epoch of reionization, when the universe was less than 0.8 Gyr old, have been made possible by large and dedicated community efforts that have identified a sample of more than 1000 galaxy candidates at $z > 7$, spanning a large range in luminosities. This progress is thanks to a variety of surveys, including ultra-deep observations in blank fields with the HUDF09 and HUDF12 campaigns (Bouwens et al. 2011; Dunlop et al. 2013; Illingworth et al. 2013), the use of cluster-scale lensing to probe intrinsically faint objects (e.g., Yue et al. 2014; Atek et al. 2015b), the large area, panchromatic CANDELS survey (Grogin et al. 2011; Koekemoer et al. 2011), and wide-field random-pointing surveys to identify more luminous but rarer objects at the bright end of the LF (e.g., Trenti et al. 2011; Bradley et al. 2012; Schmidt et al. 2014).

The picture emerging from these observations is that the number density of galaxies decreases with increasing redshift, while the LF remains consistent with a Schechter (1976) form,

$\Phi(L) = \Phi^*(L/L^*)^\alpha \exp(-L/L^*)/L^*$ out to $z \sim 8$, albeit the uncertainty and co-variance in the best fitting parameters grows at high redshift (Schmidt et al. 2014; Bouwens et al. 2015a). However, other studies find either an equally good fit with a double power law at $z \sim 8$ (Finkelstein et al. 2015a), or a preference for the double power law at $z \sim 7$ (Bowler et al. 2014, 2015; but note that the $z \sim 7$ investigations are based on ground, rather than space observations). Irrespective of the form of the LF at the bright-end, whose evolution might be linked to changing astrophysical conditions of high-redshift galaxies, such as reduced feedback at early times (Somerville et al. 2008), it is clear that overall, the observed population of galaxies ($M_{AB} \lesssim -17$) at $z > 7$ does not produce sufficient photons to ionize the universe. This is consistent both with theoretical modeling (Trenti et al. 2010; Raičević et al. 2011; Robertson et al. 2015) and indirect probes, such as the lack of detections of gamma-ray burst host galaxies (Trenti et al. 2012; Tanvir et al. 2012), which suggest that even the deepest HST observations are seeing only the tip of the iceberg of the population of star-forming galaxies. Therefore, it is likely that a very faint population of unseen dwarf galaxies at $z \gtrsim 6$ is the main contributor to the UV luminosity density and in turn to the ionizing photon budget (e.g., see Alvarez et al. 2012). Indirect observational support for extrapolation of the UV LF is provided by the detection of ultrafaint galaxies behind gravitational lenses at $z \sim 2$ (Alavi et al. 2014). However, see Giallongo et al. (2015) and Madau & Haardt (2015) for a discussion of the potential contribution of active galactic nuclei to reionization.

HST is both wavelength and aperture limited to observations at $z \lesssim 11$, but future progress to characterize the properties of the UV LF at high redshift will be boosted by the upcoming

James Webb Space Telescope (*JWST*; Gardner et al. 2006) and *Wide-field Infrared Survey Telescope* (*WFIRST*; Spergel et al. 2015). These observatories are expected to extend the frontier of galaxy detection to before the epoch of reionization at $z > 10$, when the first generation of galaxies were being assembled and hydrogen in the universe was predominantly neutral.

Motivated by this upcoming improvement in discovery capabilities of high-redshift galaxies, we aim here at predicting the UV LFs at $z > 10$. For this we use a simple, yet successful, semi-analytic framework introduced previously by Trenti et al. (2010, 2015) and Tacchella et al. (2013) that assumes that the main driver of the evolution of the galaxy UV LF is the growth and hierarchical assembly of dark matter (DM) halos. While this simple approach does not have the power of numerical hydrodynamic simulations (e.g., Genel et al. 2014; Furlong et al. 2015) to make predictions about galaxy properties such as morphology, and the role played by satellite galaxies, it provides a simple, robust and empirically calibrated method to make predictions for the evolution of global galaxy properties. In particular, the simplicity of our approach avoids the degeneracies of large multi-parameter numerical simulations, and allows us to calibrate at one reference redshift the complex physics that regulates the conversion of baryons into stars, and then to focus on investigating how DM halo assembly drives redshift evolution under the assumption that star formation efficiency is redshift independent at fixed halo mass.

The key assumptions of the framework are that, for star-forming galaxies, the stellar mass content of a DM halo depends on the halo mass but not on redshift, and that the stellar mass built up has a characteristic timescale given by the halo assembly time. The mass-dependent efficiency of converting halo mass into stellar mass is calibrated empirically at one redshift, where good observational constraints are available (e.g., $z = 4-5$), and then applied to predict the redshift evolution of the UV LF by combining evolution in the DM halo mass function (HMF) and halo assembly time with stellar population synthesis modeling. Since halos assembled more rapidly in the past, the SFR was higher at high redshift, explaining the general brightening of the luminosity versus halo mass relation found by abundance matching studies (e.g., Cooray & Milosavljević 2005).

This simple strategy has been remarkably successful in describing the UV LF evolution with redshift (Tacchella et al. 2013), with results similar to those obtained by other studies based on matching DM halos to galaxy luminosity (e.g., Lacey et al. 2011; Muñoz 2012; Jaacks et al. 2013; Behroozi & Silk 2014; Dayal et al. 2014; Mashian et al. 2015). However, the earlier implementations had two limitations. First, the model took into account only the star formation happening in the second half of the halo assembly history: the time the halo took to grow from $M_h/2$ to M_h . In addition, the calibration of the star formation efficiency was not guaranteed to be internally self-consistent with the halo assembly history of a DM halo over redshift. Given that the UV luminosity of a galaxy is only weakly sensitive to the star formation history of stellar populations with ages greater than a few hundred Myr (Madau & Dickinson 2014), these limitations had very little impact at $z \lesssim 8$. However the situation is potentially different at early times, when the assembly time goes below 100 Myr. Thus, the previous version of the model could not be trusted to formulate predictions at $z \geq 10$.

In this paper, we develop and present an improved and self-consistent model that describes the full star formation history of a DM halo, and we verify that (1) it continues to describe well the UV LF at $z \lesssim 8$; (2) it successfully reproduces the latest observations at $z \gtrsim 8$, providing an important validation of the approach introduced before such observations were available. In addition, we make detailed predictions for future surveys at $z \geq 10$ with the *JWST* and the *WFIRST*, and we discuss the implications of our results in the context of the ionizing photon budget.

This paper is organized as follows: in Section 2 we introduce our new model and its calibration; Section 3 describes our model results, and our predictions for future surveys; and we summarize and conclude in Section 4. All magnitudes are AB magnitudes and we use the Planck Collaboration et al. (2015) cosmology, with $\Omega_m = 0.315$, $\Omega_\Lambda = 0.685$, $\Omega_b = 0.0490$, $h = 0.6731$, $\sigma_8 = 0.829$, and $n_s = 0.9655$.

2. MODEL DESCRIPTION

Our model considers the growth of DM halos to be the most important driving force in the growth of galaxies, and aims at predicting the evolution of the UV LF with a minimal number of assumptions. We thus assume that the SFR is proportional to: (1) the halo mass, through a mass-dependent but redshift-independent efficiency $\varepsilon(M_h)$, which is the ratio of the stellar mass formed during the halo assembly time $t_a(M_h, z)$ to the final halo mass; and (2) to the inverse of the halo assembly time. The halo assembly time, $t_a(M_h, z)$, of a halo of mass M_h observed at redshift z is the lookback time at which the progenitor halo had mass $M_h/2$ (Lacey & Cole 1993), and decreases at higher redshift. This implies that DM halos at fixed mass host star-forming galaxies with stellar mass that is independent of redshift, but with stellar populations that are younger at higher redshift.

Since the HMFs and assembly times are fully defined by the cosmological model parameters, these assumptions allow us to calibrate our model at one redshift to derive $\varepsilon(M_h)$, and then construct predictions for the galaxy LF at all other redshifts from the DM HMF and the halo assembly time.

2.1. Star Formation Prescription

The UV luminosity of a galaxy most strongly depends on its youngest stars, while stellar populations older than a few hundred Myr contribute little. However, at $z \lesssim 8$, the halo assembly time is less than 100 Myr, so to predict the UV luminosity at high redshift it is necessary to consider periods of star formation before the halo assembly time.

To include multiple epochs of star formation, as a halo grows in DM mass, we define the star formation history of a halo as a linear combination of constant bursts normalized by the length of each burst. Thus we define the SFR for a halo with mass M_h in each epoch between times t_i and t_{i+1} as

$$\text{SFR}(t_i, t_{i+1}, M_h) = M_h \times \frac{\varepsilon(M_h/2^i)}{2^i(t_{i+1} - t_i)} \quad (1)$$

where we define t_0 as the lookback time for a halo observed at redshift z_{obs} and $t_{i>0} = t_a(M_h/2^{i-1}, z_{i-1})$, where $t_a(z_a)$ is the halo assembly time (redshift). We similarly define $z_0 = z_{\text{obs}}$ and $z_{i>0} = z_a(M_h/2^{i-1}, z_{i-1})$. We calculate the halo assembly time as defined by Lacey & Cole (1993) in the extended Press–

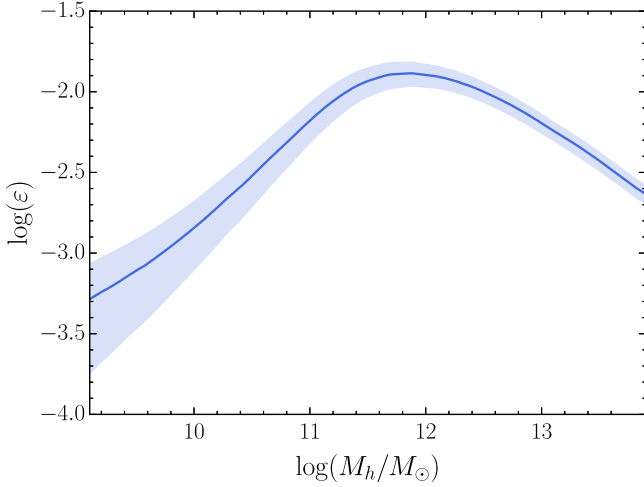


Figure 1. Efficiency of star formation, the ratio of stellar mass formed during the halo assembly time to halo mass, see Equation (2), derived at the calibration redshift $z \sim 5$, as described in Section 2.3. The shaded region shows 1σ confidence range.

Schechter formalism (Bond et al. 1991) using an ellipsoidal collapse model (Sheth et al. 2001; Giocoli et al. 2007). We use the median of the probability distribution of assembly times for each halo. While this assumption does not take into account variations in the luminosity of individual galaxies, there is a minimal effect on the global LF from neglecting scatter in halo assembly times, as demonstrated by Tacchella et al. (2013).

We define the redshift-independent efficiency of star formation, $\varepsilon(M_h)$, as the ratio of the stellar mass formed during the halo assembly time to the final halo mass. Thus, to make predictions, we only need to calibrate $\varepsilon(M_h)$ at one redshift (see Section 2.3 and Figure 1) and can use the derived $\varepsilon(M_h)$ for all further predictions.

Figure 2 shows the star formation history of halos of fixed final mass $M_h = 10^{11} M_\odot$ observed at $z_0 = 2, 6$ and 10 , calculated using the SFR in Equation (1). As $\varepsilon(M_h)$ is redshift independent, these halos will also have identical stellar masses at their observed redshifts. The SFR shown in Figure 2 increases in each epoch as the halo grows from $M_h/64$ to M_h , because $\varepsilon(M_h)$ decreases with decreasing halo mass (for $M_h \lesssim 10^{12} M_\odot$) more rapidly than the shortening of the halo assembly times as the lookback time grows. This behavior of our model is fully consistent with strong evidence of rising star formation histories with redshift from both numerical simulations and observations (Finlator et al. 2011; Papovich et al. 2011; Jaacks et al. 2012; Behroozi et al. 2013b; Lee et al. 2014). Thus the greatest contribution to the stellar mass is during the halo assembly time as the halo grows from $M_h/2$ to M_h . This figure also illustrates how the short halo assembly times at high redshift require a considerably higher SFR to form the same final stellar mass.

We include the contribution from star formation in successively smaller halo progenitors by summing the terms from Equation (1). The sum is truncated when the progenitor halo mass is below the cooling threshold, i.e., $\varepsilon(M_h \lesssim 10^8 M_\odot) = 0$.

Thus we can derive the stellar mass as

$$M_*(M_h) = M_h \times \sum_{i=0}^{\infty} \frac{\varepsilon(M_h/2^i)}{2^i}, \quad (2)$$

which is redshift independent.

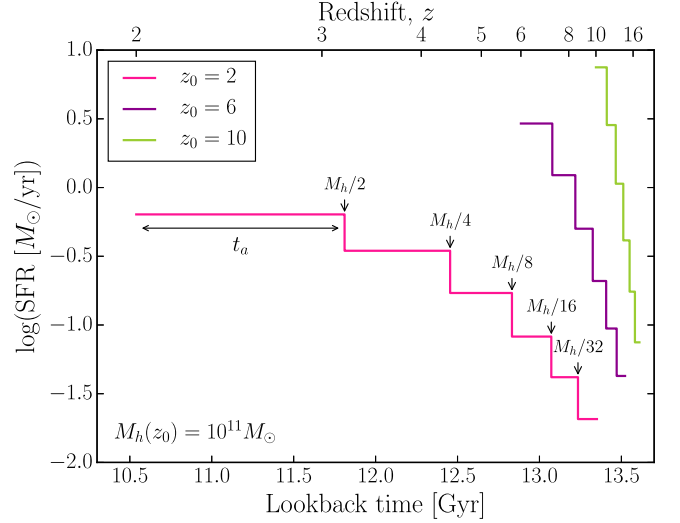


Figure 2. Star formation history, as described by Equation (1) in our model, of a halo of fixed mass $10^{11} M_\odot$, if observed at $z_0 = 2, 6$, or $z_0 = 10$. We label the halo assembly time for the $z_0 = 2$ halo and the mass of the halo progenitor at the start of each constant star formation epoch.

To compute the UV luminosity of a halo we populate every halo with a galaxy with a stellar population based on the simple stellar population (SSP) models of Bruzual & Charlot (2003). We assume a Salpeter initial mass function (IMF) between $0.1 M_\odot$ and $100 M_\odot$, as low mass stars do not contribute much to UV luminosity, and constant stellar metallicity $Z = 0.02 Z_\odot$. We neglect redshift evolution in metallicity as the UV luminosity does not depend strongly on metallicity under the assumption that current and future *HST*/*JWST* surveys detect galaxies living in relatively massive DM halos ($M_h \gtrsim 10^9 M_\odot$) where multiple generations of early star formation began enriching the gas at $z \gtrsim 20$ – 40 (Bromm et al. 2009; Trenti & Stiavelli 2009; Smith et al. 2015), so that by $z \lesssim 16$ the metallicity has risen to $Z \gtrsim 0.01 Z_\odot$. We define $\ell_{bc}(t)$ as the luminosity at 1500 \AA of an SSP of mass $1 M_\odot$ and age t . The total UV luminosity of a galaxy observed at redshift z is obtained by integrating over the SFR (Equation (1)) and SSP luminosity in each epoch of star formation:

$$L(M_h, z) = \sum_{i=0}^{\infty} \int_{t_i}^{t_{i+1}} \text{SFR}(t_i, t_{i+1}, M_h) \ell_{bc}(t) dt \quad (3)$$

where t_i are the halo assembly times for the half-mass progenitors defined above.

2.2. Dust Extinction

The observed UV luminosity is significantly attenuation by dust extinction, particularly at $z \lesssim 4$. Thus, we include dust extinction in our model, following closely the procedure adopted in observations of Lyman-break galaxies. We assume a spectrum modeled as $f_\lambda \sim \lambda^\beta$, and extinction $A_{uv} = 4.43 + 1.99\beta$ (Meurer et al. 1999). Following Trenti et al. (2015) and Tacchella et al. (2013) we model the

observations of β by Bouwens et al. (2014) as

$$\langle \beta(z, M_{\text{uv}}) \rangle = \begin{cases} (\beta_{M_0}(z) - c) \exp \left[-\frac{d\beta/dM_0(z)[M_{\text{uv}} - M_0]}{\beta_{M_0}(z) - c} \right] + c & M_{\text{uv}} \geq M_0 \\ \frac{d\beta}{dM_0}(z)[M_{\text{uv}} - M_0] + \beta_{M_0}(z) & M_{\text{uv}} < M_0 \end{cases} \quad (4)$$

where $c = -2.33$, $M_0 = -19.5$, and the values of β_{M_0} and $d\beta/dM_0$ are taken from Table 3 in Bouwens et al. (2014) and linearly extrapolated to lower and higher redshifts. Using this linear plus exponential model for β we fit the Bouwens et al. (2014) observations well—which show evidence for a curved relation between β and M_{uv} . The exponential fit at faint magnitudes avoids unphysical negative dust corrections.

We assume a Gaussian distribution for β at each M_{uv} value (with dispersion $\sigma_\beta = 0.34$), giving the average extinction $\langle A_{\text{uv}} \rangle = 4.43 + 0.79 \ln(10) \sigma_\beta^2 + 1.99 \langle \beta \rangle$. We use this average extinction to calculate observed UV fluxes in our model.

2.3. Calibration

We calibrate the model by finding $\varepsilon(M_h)$ such that $L(M_h, z_c) = L^{\text{obs}}(M_h, z_c)$, where z_c is the calibration redshift. To find $L^{\text{obs}}(M_h, z_c)$ we derive an empirical relation between observed luminosity and halo mass by performing abundance matching (Mo et al. 1999) between the HMF, $n(M_h, z)$, and observed LF, $\Phi(L, z)$, assuming every halo hosts one galaxy:

$$\int_{M_h}^{\infty} n(M'_h, z) dM'_h = \int_{L^{\text{obs}}}^{\infty} \Phi(L', z) dL' \quad (5)$$

where we use the Sheth & Tormen (1999) HMF and calibrate at $z \sim 5$ using the Bouwens et al. (2015a) LF. We calibrate over the halo mass range 10^7 – $10^{14} M_\odot$, and therefore extrapolate the LF over $M_{\text{uv}} = -7.5$ to $M_{\text{uv}} = -24.5$ to perform the abundance matching.

We can then calculate the UV luminosity as a function of halo mass at any redshift using Equation (3). The largest contribution to the uncertainty in the predicted luminosity is from the uncertainty in the calibration LF; we plot 1σ confidence bounds from this uncertainty in the plots that follow.

3. RESULTS

3.1. Stellar Masses and Ages

Figure 1 shows the efficiency of star formation as a function of halo mass, $\varepsilon(M_h)$, derived at the calibration redshift $z \sim 5$. Halos with mass $M_h \sim 10^{11}$ – $10^{12} M_\odot$ have the highest star formation efficiencies, which is consistent with the consensus picture that low-mass halos have lower efficiencies because of supernova feedback, while star formation in higher mass halos is affected by strong negative active galactic nucleus feedback. Our derived $\varepsilon(M_h)$ is consistent, to first approximation, with the results obtained by Behroozi et al. (2013a) who calculated the ratio of stellar to halo mass at different redshifts ($z \leq 4$) by

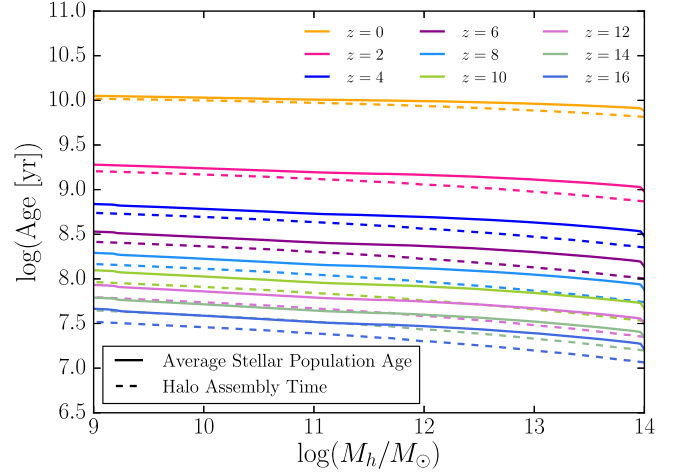


Figure 3. Average stellar population age in our model and halo assembly time (Lacey & Cole 1993) as a function of halo mass and observed redshift.

applying abundance matching techniques between the stellar mass function and HMF. We note that $\varepsilon(M_h)$ is redshift independent in our model, but that the evolution of the HMF and halo assembly times allows us to predict the UV luminosity using Equation (3), via our rising star formation history (see Figure 2). Thus we find that with our simple redshift-independent $\varepsilon(M_h)$ the star formation at high redshift proceeds much more rapidly than at low redshift.

With Equation (1), we can derive the model predictions for average galaxy ages as a function of halo mass and redshift, which are shown in Figure 3. Galaxy ages are challenging to constrain observationally, due to the degeneracy between dust extinction, age and metallicity in spectral energy distribution (SED) fitting, which are especially severe at $z \gtrsim 4$ where *HST* covers only the rest-frame UV wavelengths. Our modeling results are consistent with the picture emerging from multiple recent studies that combine *HST* and *Spitzer*/IRAC data and find that the majority of $z \gtrsim 4$ galaxies have old stellar populations (>100 Myr) and relatively low specific SFR ($\sim 100 M_\odot \text{ yr}^{-1}$; Oesch et al. 2013b; Straatman et al. 2014; see also González et al. 2011 for earlier studies reaching the same conclusion but without accounting for nebular emission lines). However, other independent studies based mostly on overlapping data sets reached the different conclusion that $z \gtrsim 4$ galaxies have a high chance of being young systems (ages <50 Myr) with high specific SFRs (de Barros et al. 2014; Finkelstein et al. 2015b), suggesting that improved observational constraints are needed to evaluate the fidelity of stellar ages predicted by our model.

The contributions to the total UV luminosity at fixed halo mass from two epochs of star formation as a function of observed redshift are shown in Figure 4. We find that the earlier epoch of star formation, as the halo grows from $M_h/4$ to $M_h/2$ adds a negligible contribution to the total UV luminosity, suggesting that recent star formation is the most important contribution to the UV luminosity even at $z > 10$ when the assembly time is short. This is consistent with recent clustering studies (Barone-Nugent et al. 2014) that found a high duty cycle for galaxies at $z \gtrsim 6.5$, possibly due to their bright, young stellar populations, and with the prediction of smoothly rising SFRs from numerical simulations (Finlator et al. 2011) and observations (Papovich et al. 2011; Behroozi et al. 2013b; Lee

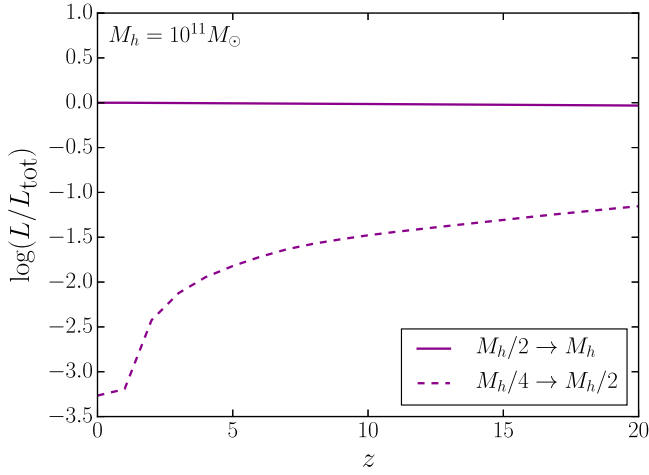


Figure 4. Contribution to predicted UV luminosity from the 0th and 1st order terms in Equation (3) as a function of observed redshift for fixed halo mass, $M_h = 10^{11} M_\odot$. The contribution from star formation during the halo assembly time (solid) dominates, with the contribution from the earlier star formation epoch (dashed) increasing with redshift. The SFR is constant in both epochs.

et al. 2014). Thus, we only consider the contribution from the first two terms in Equation (3) in predicting the UV luminosity.

The stellar mass density as a function of redshift, obtained by integrating our model stellar mass functions (using Equation (2) to derive stellar masses) to a stellar mass limit of $M_* > 10^8 M_\odot$ is shown in Figure 5. We find good agreement with observations (Pérez-González et al. 2008; Stark et al. 2013; Tomczak et al. 2014; Grazian et al. 2015; Song et al. 2015) taking into account the scatter in observations at $z < 2$, and the uncertain but potentially significant contribution of massive quiescent galaxies at low redshift.

3.2. LFs and SFR Density

The predicted observed $L(M_h, z)$ for a range of redshifts is plotted in Figure 6. Our model naturally provides redshift evolution of $L(M_h, z)$ through the evolution of the HMF and halo assembly times. Dust extinction (Section 2.2) significantly affects the high mass end of the relation at low redshift. At $z > 10$ we see the high mass end does not evolve much with redshift, motivating the model of Mashian et al. (2015) that uses an empirical redshift-independent $L(M_h)$. However, there is significant evolution at lower mass, which comprises the greatest contribution to the photon budget available to reionization because of the steep faint-end slope.

The model UV LFs at $z \leq 7$ and $z > 7$ are shown in Figures 7 and 8 respectively. The model is remarkably consistent with the observed data (Arnouts et al. 2005; Oesch et al. 2010, 2013a, 2014; Alavi et al. 2014; Bowler et al. 2015; Bouwens et al. 2015a, 2015b; Finkelstein et al. 2015a), which is expected due to the success of our previous implementation of this class of models (Trenti et al. 2010, 2015; Tacchella et al. 2013). We find the model marginally overpredicts the bright end of the LF at $z \sim 0.3$ and $z \sim 2$ by ~ 0.1 dex due to the difficulty in modeling dust extinction at these redshifts, but the observations are still within 2σ of our model. In particular, our model predicts a steepening of the faint-end slope at higher redshifts, consistent with the observed trend (Bouwens et al. 2015a).

At $z > 7$, the model describes the most recent observed data (Oesch et al. 2013a, 2014; Bouwens et al. 2015a, 2015b;

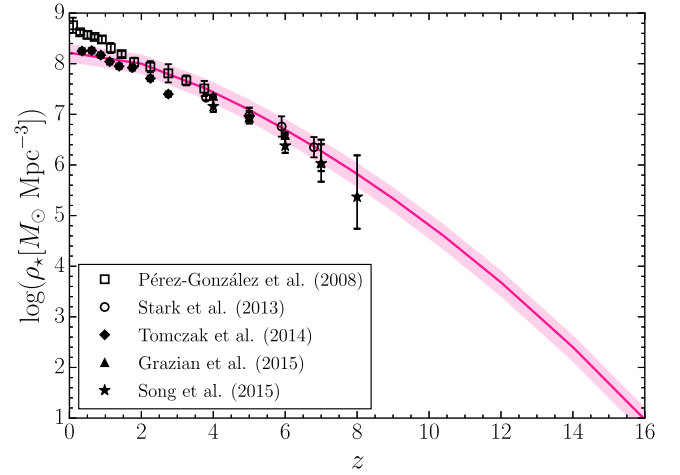


Figure 5. Stellar mass density (ρ_*) obtained by integrating the stellar mass function derived from our model using Equation (2) to a stellar mass limit of $M_* > 10^8 M_\odot$. We plot observations from Pérez-González et al. (2008); Stark et al. (2013); Tomczak et al. (2014); Grazian et al. (2015); Song et al. (2015). Shaded regions show the 1σ confidence range.

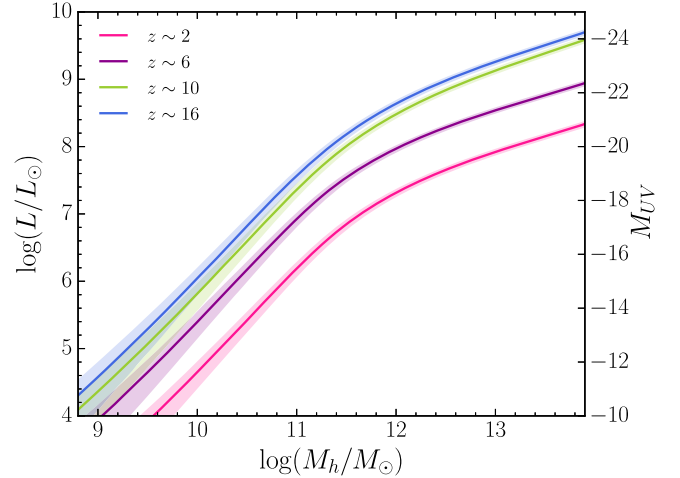


Figure 6. Relationship between observed galaxy luminosity and halo mass as a function of redshift, $L(M_h, z)$, plotted at $z = 2, 6, 8$, and 16 . Shaded regions show the 1σ confidence range.

Finkelstein et al. 2015a) well, validating our simple approach. Our model predicts the trend of steepening faint-end slope to continue at $z > 10$, and number densities to drop rapidly. With *JWST* capabilities, except in an extremely wide-field survey, the UV LF at $z > 10$ will be observed as a steep power-law function. This is agreement with the semi-analytic results of Behroozi & Silk (2014) who find a steepening power law slope at faint magnitudes and a significant drop in number densities.

For comparison, we also plot in Figure 9 the model LFs obtained by using the Finkelstein et al. (2015a) $z \sim 5$ LF for the calibration. The Finkelstein et al. (2015a) $z \sim 5$ LF has a lower value of M^* than that of Bouwens et al. (2015a), so it is not surprising that this calibration produces slightly lower number densities compared to our reference using the Bouwens et al. (2015a) LF. However, Figure 9 clearly shows that the overall evolution trends are unchanged and that the two calibrations produce LFs that are consistent within one standard deviation both between themselves and the observed data over $0 \lesssim z \lesssim 10$.

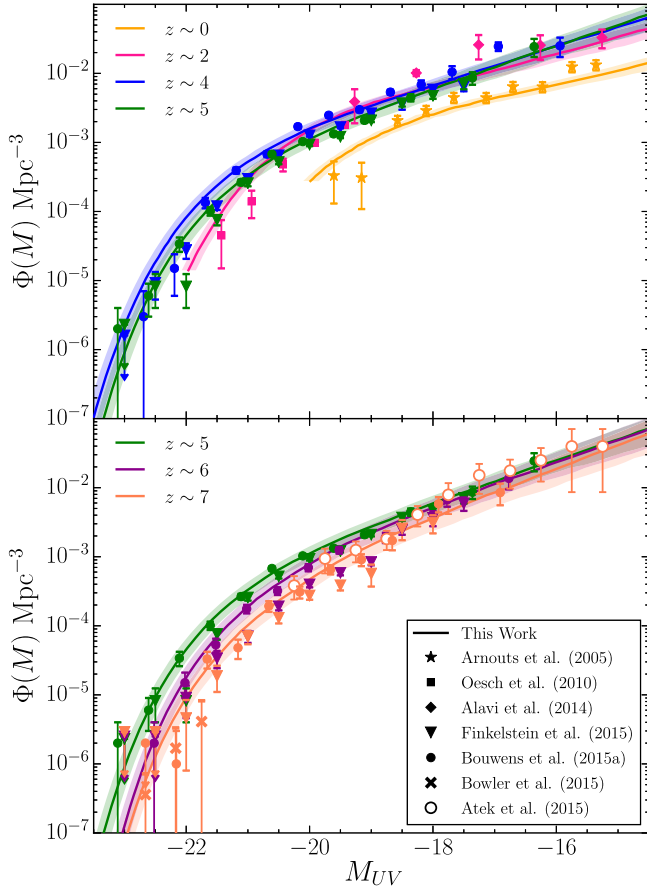


Figure 7. Predicted UV LFs at low (upper) and intermediate (lower) redshift. We show the LFs using the calibration (see Section 2.3) at $z \sim 5$ from Bouwens et al. (2015a), with Planck 2015 cosmology (Planck Collaboration et al. 2015). Points show the binned UV LFs and upper limits from Arnouts et al. (2005), Alavi et al. (2014), Oesch et al. (2010), Bouwens et al. (2015a), Finkelstein et al. (2015a), Bowler et al. (2015), Atek et al. (2015a). We note that the data from Atek et al. (2015a) were made public after our model was submitted and illustrates the consistency of our model with observations even at very low luminosity. Shaded regions show the 1σ confidence range.

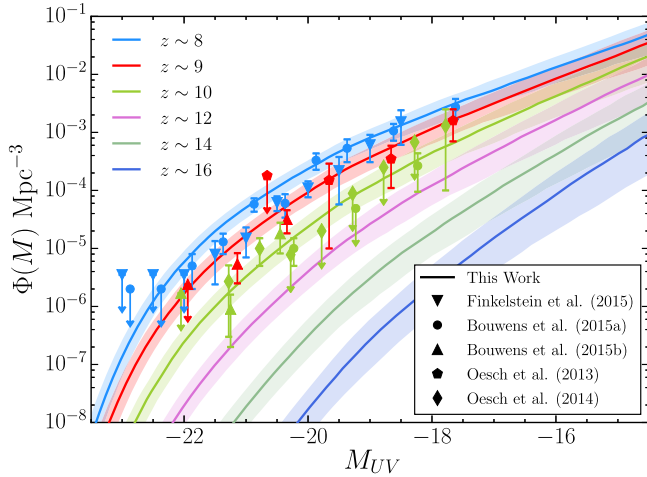


Figure 8. Predicted UV LFs at high redshift. We show the LFs using the calibration (see Section 2.3) at $z \sim 5$ from Bouwens et al. (2015a), with Planck 2015 cosmology (Planck Collaboration et al. 2015). Points show the binned UV and upper limits LFs from Oesch et al. (2013a, 2014), Finkelstein et al. (2015a), Bouwens et al. (2015a, 2015b). Shaded regions show the 1σ confidence range.

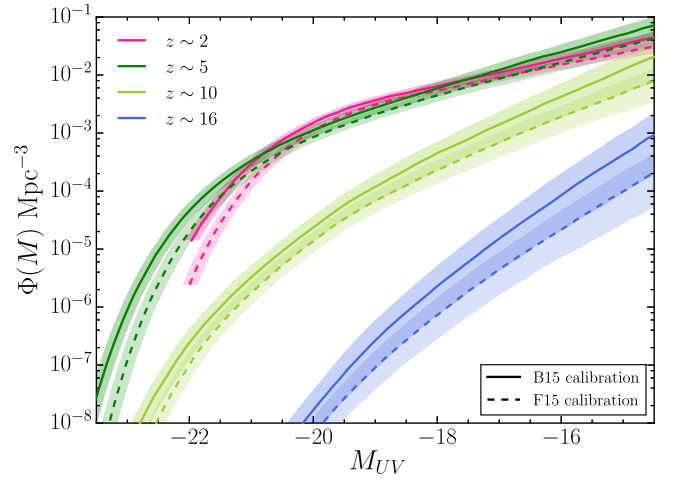


Figure 9. Predicted LFs at redshifts $z \sim 2, 5, 10, 16$ obtained by calibrating (see Section 2.3) our model with the Finkelstein et al. (2015a) LF at $z \sim 5$ (F15, dashed), compared to our reference calibration using the Bouwens et al. (2015a) LF at $z \sim 5$ (B15, solid). Shaded regions show the 1σ confidence range, highlighting that within the uncertainty of the calibrations, the two approaches yield consistent results.

Table 1
Best-fit Schechter Parameters for Model LFs

Redshift	α	M^*	$\log(\Phi^* [\text{mag}^{-1} \text{Mpc}^{-3}])$
$z \sim 0$	-1.68 ± 0.09	-19.9 ± 0.1	$-2.97^{+0.08}_{-0.07}$
$z \sim 2$	-1.46 ± 0.09	-20.3 ± 0.1	$-2.52^{+0.09}_{-0.07}$
$z \sim 4$	-1.64 ± 0.11	-21.2 ± 0.2	$-2.93^{+0.19}_{-0.13}$
$z \sim 5$	-1.75 ± 0.13	-21.2 ± 0.2	$-3.12^{+0.24}_{-0.15}$
$z \sim 6$	-1.83 ± 0.15	-20.9 ± 0.2	$-3.19^{+0.25}_{-0.16}$
$z \sim 7$	-1.95 ± 0.17	-21.0 ± 0.2	$-3.48^{+0.32}_{-0.18}$
$z \sim 8$	-2.10 ± 0.20	-21.3 ± 0.4	$-4.03^{+0.72}_{-0.26}$
$z \sim 9$	-2.26 ± 0.22	-21.2 ± 0.4	$-4.50^{+1.36}_{-0.29}$
$z \sim 10$	-2.47 ± 0.26	-21.1 ± 0.5	-5.12 ± 0.34
$z \sim 12$	-2.74 ± 0.30	-21.0 ± 0.5	-5.94 ± 0.38
$z \sim 14$	-3.11 ± 0.38	-20.9 ± 0.5	-7.05 ± 0.45
$z \sim 16$	-3.51 ± 0.46	-20.7 ± 0.6	-8.25 ± 0.51

Note. Fit is performed between $M_{\text{AB}} = -17.5$ and $M_{\text{AB}} = -22.5$

The best-fit Schechter (1976) function parameters for our LFs are shown in Table 1. The best-fit parameters are in good agreement with observations (Oesch et al. 2010; Schmidt et al. 2014; Bouwens et al. 2015a; Bowler et al. 2015; Finkelstein et al. 2015a) given the large degeneracies in Schechter function parameters. Encouragingly, we find the evolution of the derived Schechter parameters is in excellent agreement with the observed evolution (Bouwens et al. 2015a; Bowler et al. 2015): we find $d\alpha/dz \sim -0.1$, $dM^*/dz \sim 0.1$, and $d \log(\Phi^*)/dz \sim -0.3$ between $z \sim 4$ and $z \sim 8$. We find the evolution of α and Φ^* between $z \sim 8$ and $z \sim 16$ to be more dramatic: $d\alpha/dz \sim -0.2$, $dM^*/dz \sim 0.1$, and $d \log(\Phi^*)/dz \sim -0.5$, consistent with the rapid evolution of $\sim 10^{10} M_{\odot}$ halos in the DM HMF at these redshifts.

Figure 10 shows the luminosity density and cosmic SFR density as a function of redshift. We calculate the luminosity density by integrating our model LFs down to a magnitude limit. We choose two fiducial limits of $M_{\text{lim}} = -17$ (just fainter than current observational limits) and $M_{\text{lim}} = -12$ (the theoretical mass limit for halos to cool). We calculate the SFR density, $\dot{\rho}_*$ using the empirical relation from Madau et al.

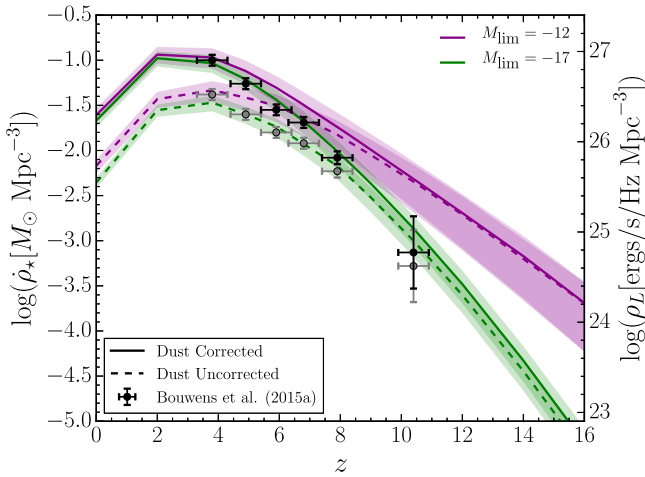


Figure 10. Luminosity density (ρ_L) and cosmic SFR density (ρ_*) as functions of redshift, derived by integrating the model UV LFs to magnitude limits of $M_{AB} = -17$ (green lines) and $M_{AB} = -12$ (purple lines). The dust corrected SFR densities for the two magnitude limits are shown as solid lines, dust uncorrected SFR densities are shown as dashed lines. The observed SFR densities from Bouwens et al. (2015a) are shown in black (dust corrected) and gray (dust uncorrected). Shaded regions show the 1σ confidence range.

(1998) where $\text{SFR} [M_\odot \text{ yr}^{-1}] = 8.0 \times 10^{27} L [\text{erg s}^{-1} \text{ Hz}^{-1}]$ at 1500 \AA . We plot the densities and observations from Bouwens et al. (2015a) both with and without dust correction. At $z \leq 8$ the observations are consistent with both magnitude limits, however the observations at $z \sim 10$ suggests a significant steepening of the relation at high redshift, as do results from numerical simulations (Genel et al. 2014, though with large uncertainty), which is consistent with our model with $M_{\text{lim}} = -17$. The sample at $z \sim 10$ is limited however; more observational data at $z > 8$ are needed to confirm this result.

3.3. Forecasts for JWST and WFIRST

We use our model to make forecasts for a representative set of *JWST* NIRCAM high-redshift dropout surveys using the 5 near-IR filters. The surveys (properties summarized in Table 2) include an ultra-deep (UD) survey of 4 pointings ($\sim 40 \text{ arcmin}^2$) exposed in 200 hr per pointing; a medium-deep (MD) survey of 40 pointings exposed in 20 hr per pointing; and a wide-field (WF) survey of 400 pointings exposed in 2 hr per pointing. We assume that the surveys will split the observing time so as to reach equal depth in all five filters, and estimate the limiting magnitude for an 8σ detection (in a single filter) using the *JWST* Exposure Time Calculator. We also include the effects of gravitational lensing magnification bias from strong lensing in blank fields, which is expected to distort the brightest end of high-redshift LFs (Wyithe et al. 2011; Mason et al. 2015).

In Figure 11 we plot the predicted cumulative number counts for redshifts $8 \leq z \leq 16$ and the regions accessible to these mock *JWST* surveys, as well as the region accessible to *WFIRST* High-Latitude Survey (HLS, Spergel et al. 2015). The estimated number of dropouts are given in Table 2.

Our model predicts a significant drop in number density from $z \sim 8$ to $z \sim 10$ compared to lower redshifts (which is also seen in the observations; Bouwens et al. 2015a, 2015b). The drop continues to high redshift, thus we find that no $z \sim 16$ galaxies would be detected in our mock *JWST* surveys. To detect 1 galaxy at $z \sim 16$ in our UD survey would require ~ 40

pointings ($\sim 400 \text{ arcmin}^2$). We find that magnification bias in blank fields does not significantly affect our model even at the brightest observable magnitudes at $z > 10$. The magnification bias effect is only noticeable in the exponential part of the LF, which is within reach only at $z \sim 8$, but too weak otherwise for power laws with slope in the range -2 to -3.5 (it is exactly neutral for faint end slope $\alpha = -2$). Mason et al. (2015) showed the lensing effect was most significant for a Schechter function LF at high redshift (see also Barone-Nugent et al. 2015). Thus we expect that without significant strong lensing, i.e., using galaxy clusters as cosmic telescopes (e.g., the *Hubble Frontier Fields* Yue et al. 2014; Atek et al. 2015b; Coe et al. 2015; Ishigaki et al. 2015), $z > 15$ is beyond the reach of *JWST*.

3.4. Implications for Reionization

The timeline of cosmic reionization depends on the balance between the recombination of free electrons with protons to form neutral hydrogen atoms, and the ionization of hydrogen atoms by Lyman continuum photons emitted by young stars. The UV luminosity density (and therefore, SFR density) at a given redshift allows us to calculate the number of photons available for reionization, and is most sensitive to the faint end of the LF. We can use this to infer the timeline of reionization by calculating the ionized hydrogen fraction, $Q(z)$, as a function of redshift given the following time-dependent differential equation:

$$\dot{Q} = \frac{\dot{n}_{\text{ion}}}{\langle n_{\text{H}} \rangle} - \frac{Q}{t_{\text{rec}}} \quad (6)$$

where \dot{n}_{ion} is the comoving number density of ionizing photons, $\langle n_{\text{H}} \rangle$ is the comoving number density of hydrogen atoms, and the recombination time of the IGM (Stiavelli et al. 2004; Robertson et al. 2015 and references therein) is

$$t_{\text{rec}}(z) = [C \alpha_B(T) n_e (1+z)^3]^{-1} \quad (7)$$

where $\alpha_B(T)$ is the case B recombination (i.e., opaque IGM) coefficient for hydrogen, $n_e = (1 + Y_p/4X_p) \langle n_{\text{H}} \rangle$ is the comoving number density of electrons (assuming singly ionized He), X_p and Y_p are the primordial hydrogen and helium abundances respectively, and $C = \langle n_{\text{H}}^2 \rangle / \langle n_{\text{H}} \rangle^2$ is the “clumping factor” that accounts for inhomogeneity in the IGM.

The production rate of ionizing photons can be related to the total UV luminosity density, ρ_L as

$$\dot{n}_{\text{ion}} = f_{\text{esc}} \xi_{\text{ion}} \rho_L \quad (8)$$

where f_{esc} is the average fraction of photons that escape galaxies to affect the IGM, and ξ_{ion} is the rate of ionizing photons per unit UV luminosity, with units Hz/erg , which depends on the IMF, metallicity, age and dust content of the stellar populations. There is an equivalent relation between \dot{n}_{ion} and SFR density (Madau et al. 1999; Shull et al. 2012), which requires the same stellar population modeling.

All of the parameters involved are difficult to estimate, and may evolve with redshift as reionization progresses and the IGM evolves (Furlanetto & Oh 2005; Shull et al. 2012). In this work, we follow Schmidt et al. (2014) and use a distribution of parameters. For two limiting magnitudes ($M_{AB} = -17$, corresponding to currently observable galaxies, and $M_{AB} = -12$, corresponding to atomic cooling halos) we assume the escape

Table 2
Predicted Number Counts for Example *JWST* and *WFIRST* Surveys

Redshift	Dropout Filter	UD ($m_{\text{lim}} = 32.0$) ~40 arcmin ²	MD ($m_{\text{lim}} = 30.6$) ~400 arcmin ²	WF ($m_{\text{lim}} = 29.3$) ~4000 arcmin ²	<i>WFIRST</i> ($m_{\text{lim}} = 26.5$) ~2000 deg ²
$z \sim 8$	F115W	197^{+104}_{-92}	548^{+259}_{-225}	1335^{+595}_{-503}	61370^{+27995}_{-22029}
$z \sim 10$	F115W	30^{+21}_{-17}	52^{+33}_{-26}	102^{+64}_{-48}	1026^{+701}_{-473}
$z \sim 12$	F150W	6^{+5}_{-4}	10^{+8}_{-6}	13^{+10}_{-7}	47^{+41}_{-25}
$z \sim 14$	F150W	$0.3^{+0.4}_{-0.2}$	$0.4^{+0.4}_{-0.2}$	$0.4^{+0.4}_{-0.3}$	$0.4^{+0.4}_{-0.2}$
$z \sim 16$	F200W	0	0	0	0

Note. Limiting magnitudes for a 8σ detection estimated with the *JWST* Exposure Time Calculator and *WFIRST*HLs. The mock surveys are described in Section 3.3. These estimates include the boost from gravitational lensing magnification bias in blank fields (Wyithe et al. 2011; Mason et al. 2015).

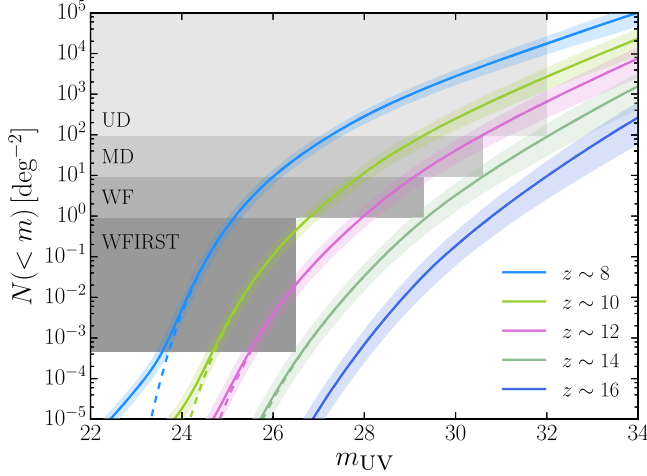


Figure 11. Predicted number counts of galaxies brighter than apparent magnitude m_{UV} (rest-frame UV) per square degree for a range of redshifts based on our model LFs. We plot the cumulative number counts including the boost from gravitational lensing magnification bias (Wyithe et al. 2011; Mason et al. 2015) as solid lines, and without the magnification bias effect (dashed lines). We plot the estimated coverage of future surveys as shaded regions: 3 mock *JWST* surveys detailed in Section 3.3 and the *WFIRST* High-latitude Survey (Spergel et al. 2015). The calculated number counts are given in Table 2.

fraction is uniformly distributed between $f_{\text{esc}} = 0.1\text{--}0.3$ (Ouchi et al. 2009), and we use a uniform distribution between $C = 1\text{--}6$ for the clumping factor. Finally we model ξ_{ion} as a log-normal distribution with mean $\log \xi_{\text{ion}} = 25.2$ and standard deviation 0.15 dex, which was the range obtained by Robertson et al. (2013) using the standard Bruzual & Charlot (2003) models and measurements of the UV spectral slope by Dunlop et al. (2012). We assume an IGM temperature of 20,000 K.

Once the reionization history, $Q(z)$, is known, an important constraint is to compare the electron scattering optical depth with that inferred from cosmic microwave background observations. The Planck Collaboration et al. (2015) reported a reionization value of $\tau = 0.066 \pm 0.012$, consistent with instantaneous reionization at $z = 8.8^{+1.2}_{-1.1}$. The optical depth as a function redshift is

$$\tau(z) = \int_0^z \sigma_T n_e (1 + z')^2 Q(z') \frac{c}{H(z')} dz' \quad (9)$$

where c is the speed of light, σ_T is the Thomson scattering cross section and $H(z)$ is the Hubble parameter.

Figure 12 shows the reionization history: the ionized hydrogen fraction as a function of redshift, obtained by solving Equation (6) with our model luminosity density, sampling the

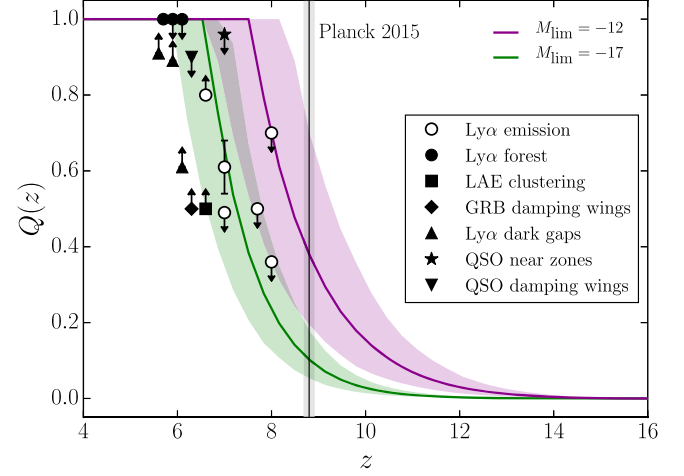


Figure 12. Fraction of ionized hydrogen as a function of redshift, obtained by solving Equation (6) with our model luminosity density. We plot our results from integrating the model UV LFs to two magnitude limits of $M_{\text{AB}} = -17$ (green) and $M_{\text{AB}} = -12$ (purple), with 1σ confidence regions as shaded regions. We also plot constraints derived from observations of: $\text{Ly}\alpha$ emission from galaxies (open circles, Ouchi et al. 2010; Faisst et al. 2014; Pentericci et al. 2014; Schenker et al. 2014; Tilvi et al. 2014); the $\text{Ly}\alpha$ forest (filled circles, Fan et al. 2006); the clustering of $\text{Ly}\alpha$ emitting galaxies (square, Ouchi et al. 2010); GRB spectra damping wings (diamond, McQuinn et al. 2008); dark gaps in the $\text{Ly}\alpha$ forest (upper triangles, McGreer et al. 2015); quasar near zones (star, Venemans et al. 2015); and quasar spectra damping wings (lower triangle, Schroeder et al. 2013). We also plot the Planck Collaboration et al. (2015) redshift of instantaneous reionization. We note that the conversion from the $\text{Ly}\alpha$ escape fraction to the global ionized hydrogen fraction is uncertain and relies on several model assumptions (Mesinger et al. 2015).

distributions of input parameters. Figure 13 shows the electron scattering optical depth as a function of redshift. For the LF magnitude limit $M_{\text{AB}} = -17$, reionization is complete ($Q = 1$) by $z_{\text{reion}} = 6.86^{+0.32}_{-0.66}$, with $\tau(z_{\text{reion}}) = 0.042^{+0.008}_{-0.002}$. For the LF magnitude limit $M_{\text{AB}} = -12$, reionization is complete by $z_{\text{reion}} = 7.84^{+0.65}_{-0.98}$, with $\tau(z_{\text{reion}}) = 0.056^{+0.007}_{-0.010}$.

The fainter magnitude limit, corresponding to atomic cooling halos of mass $M_h \sim 10^9 M_\odot$, is fully consistent with the Planck results, considering the uncertainty in the reionization model parameters. This calculation shows that ultrafaint galaxies can in principle provide enough photons to fully reionize the universe by $z \sim 6$ to match observations of the $\text{Ly}\alpha$ forest (Fan et al. 2006). Both magnitude limits are broadly consistent with a range of constraints from observations, within the reionization model uncertainty: UV luminosity densities (Finkelstein et al. 2012) for observable galaxies; quasar near zones (Venemans et al. 2015); quasar spectra damping wings (Schroeder et al. 2013); GRB spectra damping wings

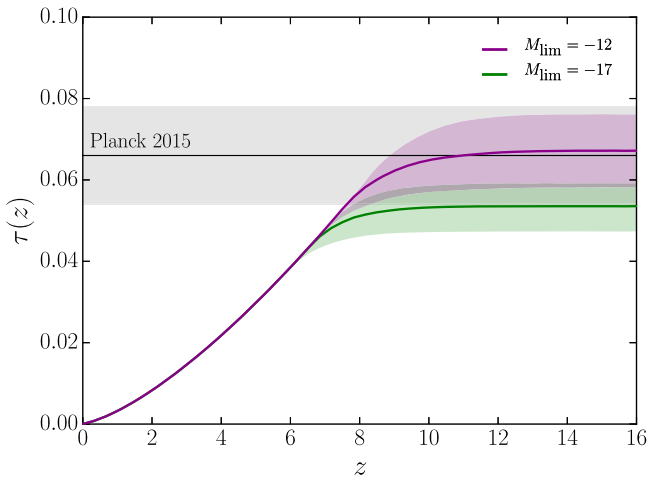


Figure 13. Electron scattering optical depth, calculated using Equation (9) from our derived $Q(z)$. We plot our results from integrating the model UV LFs to two magnitude limits of $M_{AB} = -17$ (green) and $M_{AB} = -12$ (purple), with 1σ confidence regions as shaded regions. We show the reionization optical depth value and its 1σ confidence levels from Planck Collaboration et al. (2015) in gray.

(McQuinn et al. 2008); transmission (Fan et al. 2006) and dark gaps (McGreer et al. 2015) in the Ly α forest; and the clustering of Ly α emitting galaxies (Ouchi et al. 2010).

Qualitatively, the non-negligible neutral fraction predicted by our model at $z \gtrsim 7$ is consistent with the observed high optical depth of Ly α (Ouchi et al. 2010; Treu et al. 2013; Pentericci et al. 2014; Schenker et al. 2014; Tilvi et al. 2014; Faisst et al. 2014; Schmidt et al. 2015), however the conversion from the Ly α emission fraction to the volume filling factor of ionized hydrogen is difficult and requires several assumptions (Mesinger et al. 2015). In particular, to make constraints on reionization it is generally assumed that there are no changes in galaxy and the Ly α emission line properties, which necessitates a rapid evolution of the global ionization fraction between $z \sim 6$ and $z \sim 7$. However, recent studies have shown that the rapid decline in the Ly α escape fraction at these redshifts cannot result only from the changing IGM attenuation (Mesinger et al. 2015) but could also be explained by the co-evolution of the escape fraction of ionizing photons, f_{esc} , (Dijkstra et al. 2014). Thus, the uncertainties in the ionization fraction from the Ly α optical depth shown in our plot are likely underestimated, since they do not include these systematic effects.

4. SUMMARY AND CONCLUSIONS

We have presented a simple model for the evolution of the UV LF from $0 \lesssim z \lesssim 16$, assuming that the average star formation history of galaxies is set by their halo mass and by the redshift (through the halo assembly time), so that halos of the same mass have the same stellar mass content independent of redshift. Our model builds upon previous similar implementations, but here we extended our framework to construct a self-consistent model that is capable of following the evolution of the star formation even when the halo assembly times become very short (at $z \gtrsim 10$).

Our key findings are as follows.

1. Our model UV LFs are very successful in matching observations at all redshifts where data are available ($0 \lesssim$

$z \lesssim 10$). Overall, we find that the shape of the LF is well described by a Schechter function with faint-end slope increasing with redshift. This trend continues at higher redshift, and we use the model to make predictions for LFs at $z > 10$, finding a faint-end slope $\alpha \sim -3.5$ at $z = 16$.

2. Our model reproduces the observed cosmic SFR density well, indicating a sharp decline at $z > 8$ with a magnitude limit of $M_{AB} = -17$, consistent with observed data at $z \sim 10$.
3. Compared to previous more basic models (Tacchella et al. 2013; Trenti et al. 2015), we find that the self-consistent inclusion of earlier periods of star formation does not significantly affect the total UV luminosity at a given halo mass and redshift, but it allows us to better reproduce the observed average stellar ages and stellar mass density of high-redshift galaxies.
4. Taking advantage of the ability of the model to make predictions at the earliest times, we investigate the expected galaxy detections for future ultra-deep, medium-deep and wide-field surveys with *JWST*, and *WFIRST*. We predict that $z \sim 14$ galaxies over a range of luminosities are in reach of these surveys. However, significant strong lensing magnification will be needed to push beyond $z > 15$.
5. Finally, we investigate the implications of our model for the reionization process and find that reionization is complete by $z_{\text{reion}} = 7.84^{+0.65}_{-0.98}$, under the assumption that the LF extends down to a minimum galaxy luminosity of $M_{AB} = -12$ ($M_h \sim 10^9 M_\odot$), with $\tau(z_{\text{reion}}) = 0.056^{+0.007}_{-0.010}$. Overall our model is consistent with the Planck Collaboration et al. (2015) results and with ultrafaint galaxies being the dominant sources of reionization, despite the fact that this population is currently not detected via direct imaging (but inferred indirectly through GRB host galaxy searches at $z > 6$).

This work was supported by the *HST* BoRG grants GO-12572, 12905, and 13767, and the *HST* GLASS grant GO-13459. T.T. acknowledges support by the Packard Foundation through a Packard Fellowship.

REFERENCES

- Alavi, A., Siana, B., Richard, J., et al. 2014, *ApJ*, **780**, 143
 Alvarez, M. A., Finlator, K., & Trenti, M. 2012, *ApJL*, **759**, L38
 Arnouts, S., Schiminovich, D., Ilbert, O., et al. 2005, *ApJL*, **619**, L43
 Atek, H., Richard, J., Jauzac, M., et al. 2015a, *ApJ*, in press (arXiv:1509.06764)
 Atek, H., Richard, J., Kneib, J.-P., et al. 2015b, *ApJ*, **800**, 18
 Barone-Nugent, R. L., Trenti, M., Wyithe, J. S. B., et al. 2014, *ApJ*, **793**, 17
 Barone-Nugent, R. L., Wyithe, J. S. B., Trenti, M., et al. 2015, *MNRAS*, **450**, 1224
 Behroozi, P. S., & Silk, J. 2014, *ApJ*, **799**, 32
 Behroozi, P. S., Wechsler, R. H., & Conroy, C. 2013a, *ApJL*, **762**, L31
 Behroozi, P. S., Wechsler, R. H., & Conroy, C. 2013b, *ApJ*, **770**, 57
 Bond, J. R., Cole, S., Efstathiou, G., & Kaiser, N. 1991, *ApJ*, **379**, 440
 Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2011, *ApJ*, **737**, 90
 Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2014, *ApJ*, **793**, 115
 Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2015a, *ApJ*, **803**, 34
 Bouwens, R. J., Oesch, P. A., Labbe, I., et al. 2015b, *ApJ*, submitted (arXiv:1506.01035)
 Bowler, R. A. A., Dunlop, J. S., McLure, R. J., et al. 2014, *MNRAS*, **440**, 2810
 Bowler, R. A. A., Dunlop, J. S., McLure, R. J., et al. 2015, *MNRAS*, **452**, 1817
 Bradley, L. D., Trenti, M., Oesch, P. A., et al. 2012, *ApJ*, **760**, 108
 Bromm, V., Yoshida, N., Hernquist, L., & McKee, C. F. 2009, *Natur*, **459**, 49
 Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000

- Burgarella, D., Pérez-González, P. G., Tyler, K. D., et al. 2006, *A&A*, **450**, 69
- Coe, D., Bradley, L., & Zitrin, A. 2015, *ApJ*, **800**, 84
- Cooray, A., & Milosavljević, M. 2005, *ApJL*, **627**, L89
- Dayal, P., Ferrara, A., Dunlop, J. S., & Pacucci, F. 2014, *MNRAS*, **445**, 2545
- de Barros, S., Schaerer, D., & Stark, D. P. 2014, *A&A*, **563**, A81
- Dijkstra, M., Wyithe, S., Haiman, Z., Mesinger, A., & Pentericci, L. 2014, *MNRAS*, **440**, 3309
- Dunlop, J. S., McLure, R. J., Robertson, B. E., et al. 2012, *MNRAS*, **420**, 901
- Dunlop, J. S., Rogers, A. B., McLure, R. J., et al. 2013, *MNRAS*, **432**, 3520
- Faisst, A. L., Capak, P., Carollo, C. M., Scarlata, C., & Scoville, N. 2014, *ApJ*, **788**, 87
- Fan, X., Strauss, M. A., Becker, R. H., et al. 2006, *AJ*, **132**, 117
- Finkelstein, S. L., Papovich, C., Ryan, R. E., et al. 2012, *ApJ*, **758**, 93
- Finkelstein, S. L., Ryan, R. E., Jr., Papovich, C., et al. 2015a, *ApJ*, **810**, 71
- Finkelstein, S. L., Song, M., Behroozi, P., et al. 2015b, *ApJ*, submitted (arXiv:1504.00005)
- Finlator, K., Oppenheimer, B. D., & Davé, R. 2011, *MNRAS*, **410**, 1703
- Furlanetto, S. R., & Oh, S. P. 2005, *MNRAS*, **363**, 1031
- Furlong, M., Bower, R. G., Theuns, T., et al. 2015, *MNRAS*, **450**, 4486
- Gardner, J., Mather, J., Clampin, M., et al. 2006, *SSRv*, **123**, 485
- Genel, S., Vogelsberger, M., Springel, V., et al. 2014, *MNRAS*, **445**, 175
- Giallongo, E., Grazian, A., Fiore, F., et al. 2015, *A&A*, **578**, A83
- Giocoli, C., Moreno, J., Sheth, R. K., & Tormen, G. 2007, *MNRAS*, **376**, 977
- González, V., Labbé, I., Bouwens, R. J., et al. 2011, *ApJL*, **735**, L34
- Grazian, A., Fontana, A., Santini, P., et al. 2015, *A&A*, **575**, A96
- Grogin, N. A., Kocevski, D. D., Faber, S. M., et al. 2011, *ApJS*, **197**, 35
- Illingworth, G. D., Magee, D., Oesch, P. A., et al. 2013, *ApJS*, **209**, 6
- Ishigaki, M., Kawamata, R., Ouchi, M., et al. 2015, *ApJ*, **799**, 12
- Jaacks, J., Nagamine, K., & Choi, J. H. 2012, *MNRAS*, **427**, 403
- Jaacks, J., Thompson, R., & Nagamine, K. 2013, *ApJ*, **766**, 94
- Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, *ApJS*, **197**, 36
- Lacey, C., & Cole, S. 1993, *MNRAS*, **262**, 627
- Lacey, C. G., Baugh, C. M., Frenk, C. S., & Benson, A. J. 2011, *MNRAS*, **412**, 1828
- Lee, S.-K., Ferguson, H. C., Somerville, R. S., et al. 2014, *ApJ*, **783**, 81
- Madau, P., & Dickinson, M. 2014, *ARA&A*, **52**, 415
- Madau, P., & Haardt, F. 2015, *ApJL*, in press (arXiv:1507.07678)
- Madau, P., Haardt, F., & Rees, M. J. 1999, *ApJ*, **514**, 648
- Madau, P., Pozzetti, L., & Dickinson, M. 1998, *ApJ*, **498**, 106
- Mashian, N., Oesch, P., & Loeb, A. 2015, *MNRAS*, submitted (arXiv:1507.00999)
- Mason, C. A., Treu, T., Schmidt, K. B., et al. 2015, *ApJ*, **805**, 79
- McGreer, I. D., Mesinger, A., & D'Odorico, V. 2015, *MNRAS*, **447**, 499
- McQuinn, M., Lidz, A., Zaldarriaga, M., Hernquist, L., & Dutta, S. 2008, *MNRAS*, **388**, 1101
- Mesinger, A., Aykatalp, A., Vanzella, E., et al. 2015, *MNRAS*, **446**, 566
- Meurer, G. R., Heckman, T. M., & Calzetti, D. 1999, *ApJ*, **521**, 64
- Mo, H. J., Mao, S., & White, S. D. M. 1999, *MNRAS*, **304**, 175
- Muñoz, J. A. 2012, *JCAP*, **2012**, 15
- Oesch, P. A., Bouwens, R. J., Carollo, C. M., et al. 2010, *ApJL*, **725**, L150
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2013a, *ApJ*, **773**, 75
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2014, *ApJ*, **786**, 108
- Oesch, P. A., Labbé, I., Bouwens, R. J., et al. 2013b, *ApJ*, **772**, 136
- Ouchi, M., Mobasher, B., Shimasaku, K., et al. 2009, *ApJ*, **706**, 1136
- Ouchi, M., Shimasaku, K., Furusawa, H., et al. 2010, *ApJ*, **723**, 869
- Papovich, C., Finkelstein, S. L., Ferguson, H. C., Lotz, J. M., & Giavalisco, M. 2011, *MNRAS*, **412**, 1123
- Pentericci, L., Vanzella, E., Fontana, A., et al. 2014, *ApJ*, **793**, 113
- Pérez-González, P. G., Rieke, G. H., Villar, V., et al. 2008, *ApJ*, **675**, 234
- Planck Collaboration Ade, P. A. R., Aghanim, N., et al. 2015, arXiv:1502.01589
- Raičević, M., Theuns, T., & Lacey, C. 2011, *MNRAS*, **410**, 775
- Robertson, B. E., Ellis, R. S., Furlanetto, S. R., & Dunlop, J. S. 2015, *ApJL*, **802**, L19
- Robertson, B. E., Furlanetto, S. R., Schneider, E., et al. 2013, *ApJ*, **768**, 71
- Schechter, P. L. 1976, *ApJ*, **203**, 297
- Schenker, M. A., Ellis, R. S., Konidaris, N. P., & Stark, D. P. 2014, *ApJ*, **795**, 20
- Schmidt, K. B., Treu, T., Bradač, M., & Vulcani, B. 2015, *ApJ*, submitted
- Schmidt, K. B., Treu, T., Trenti, M., et al. 2014, *ApJ*, **786**, 57
- Schroeder, J., Mesinger, A., & Haiman, Z. 2013, *MNRAS*, **428**, 3058
- Sheth, R. K., Mo, H. J., & Tormen, G. 2001, *MNRAS*, **323**, 1
- Sheth, R. K., & Tormen, G. 1999, *MNRAS*, **308**, 119
- Shull, J. M., Harness, A., Trenti, M., & Smith, B. D. 2012, *ApJ*, **747**, 100
- Smith, B. D., Wise, J. H., O'Shea, B. W., Norman, M. L., & Khochfar, S. 2015, *MNRAS*, **452**, 2822
- Somerville, R. S., Hopkins, P. F., Cox, T. J., Robertson, B. E., & Hernquist, L. 2008, *MNRAS*, **391**, 481
- Song, M., Finkelstein, S. L., Ashby, M. L. N., et al. 2015, *ApJ*, submitted (arXiv:1507.05636)
- Spergel, D., Gehrels, N., Baltay, C., et al. 2015, arXiv:1503.03757
- Stark, D. P., Schenker, M. A., Ellis, R., et al. 2013, *ApJ*, **763**, 129
- Stiavelli, M., Fall, S. M., & Panagia, N. 2004, *ApJ*, **600**, 508
- Straatman, C. M. S., Labbé, I., Spitler, L. R., et al. 2014, *ApJL*, **783**, L14
- Tacchella, S., Trenti, M., & Carollo, C. M. 2013, *ApJL*, **768**, L37
- Tanvir, N. R., Levan, A. J., Fruchter, A. S., et al. 2012, *ApJ*, **754**, 46
- Tilvi, V., Papovich, C., Finkelstein, S. L., et al. 2014, *ApJ*, **794**, 5
- Tomczak, A. R., Quadri, R. F., Tran, K.-V. H., et al. 2014, *ApJ*, **783**, 85
- Trenti, M., Bradley, L. D., Stiavelli, M., et al. 2011, *ApJL*, **727**, L39
- Trenti, M., Perna, R., & Jimenez, R. 2015, *ApJ*, **802**, 103
- Trenti, M., Perna, R., Levesque, E. M., Shull, J. M., & Stocke, J. T. 2012, *ApJL*, **749**, L38
- Trenti, M., & Stiavelli, M. 2009, *ApJ*, **694**, 879
- Trenti, M., Stiavelli, M., Bouwens, R. J., et al. 2010, *ApJ*, **714**, 202
- Treu, T., Schmidt, K. B., Trenti, M., Bradley, L. D., & Stiavelli, M. 2013, *ApJL*, **775**, L29
- Venemans, B. P., Bañados, E., Decarli, R., et al. 2015, *ApJL*, **801**, L11
- Wyithe, J. S. B., Yan, H., Windhorst, R. A., & Mao, S. 2011, *Natur*, **469**, 181
- Yue, B., Ferrara, A., Vanzella, E., & Salvaterra, R. 2014, *MNRAS*, **443**, 20