Classifying Stellar Spectra

By Edith Johnston

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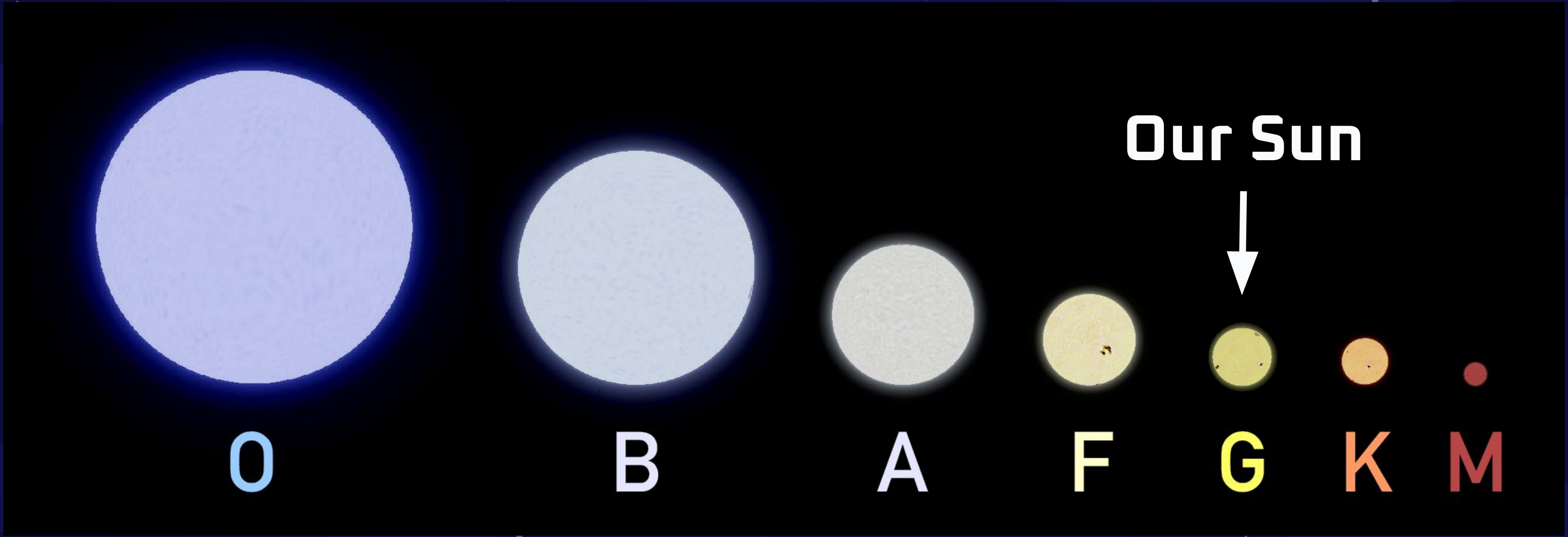
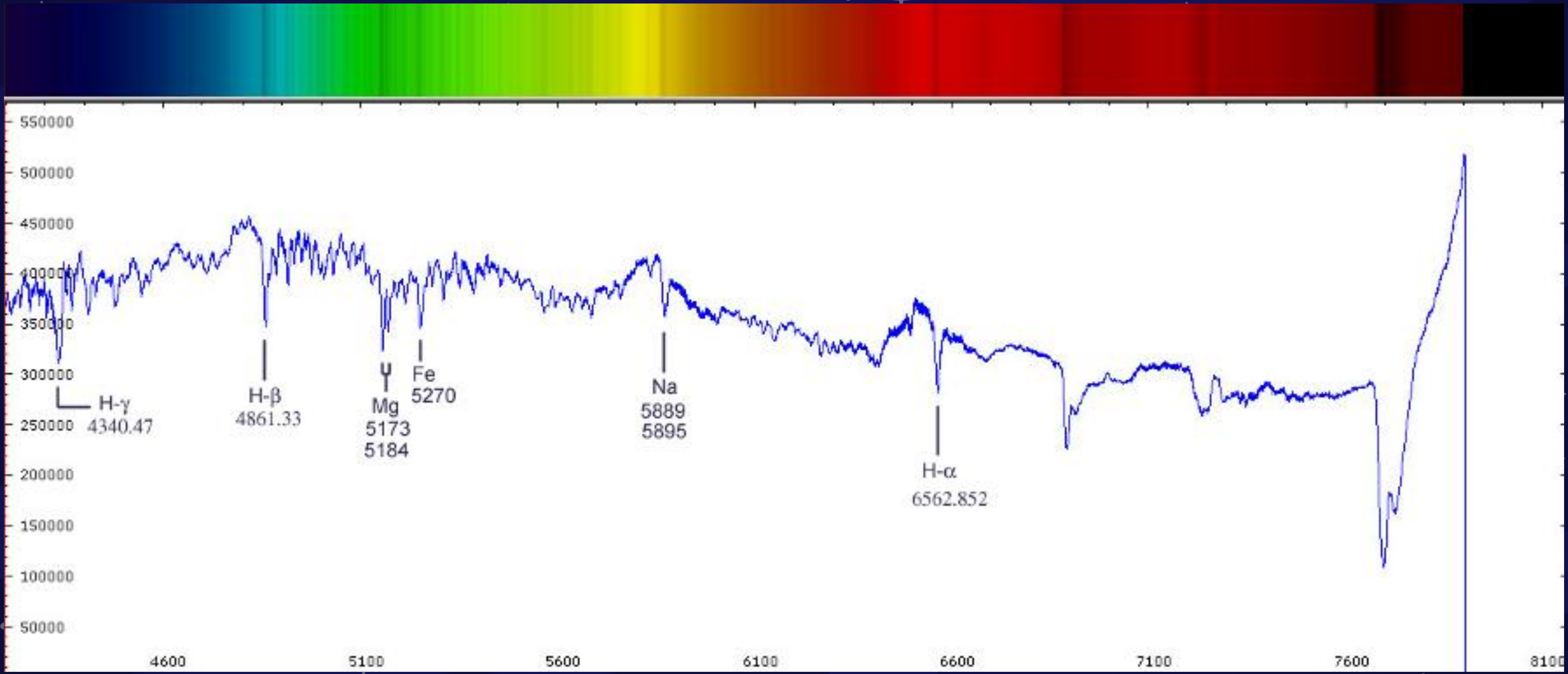
Stellar Spectral Types



Stellar Spectra



What's the Point of Classifying?

★ Different types behave differently

★ Easier to narrow down searches

- E.g. habitable planets ⇨ F, G stars

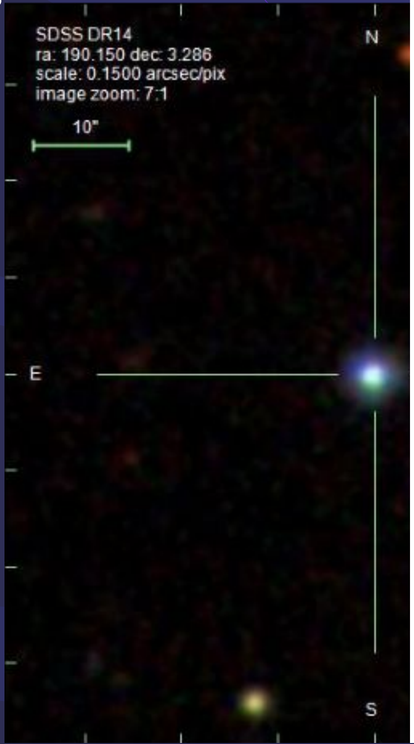
My Project Goal

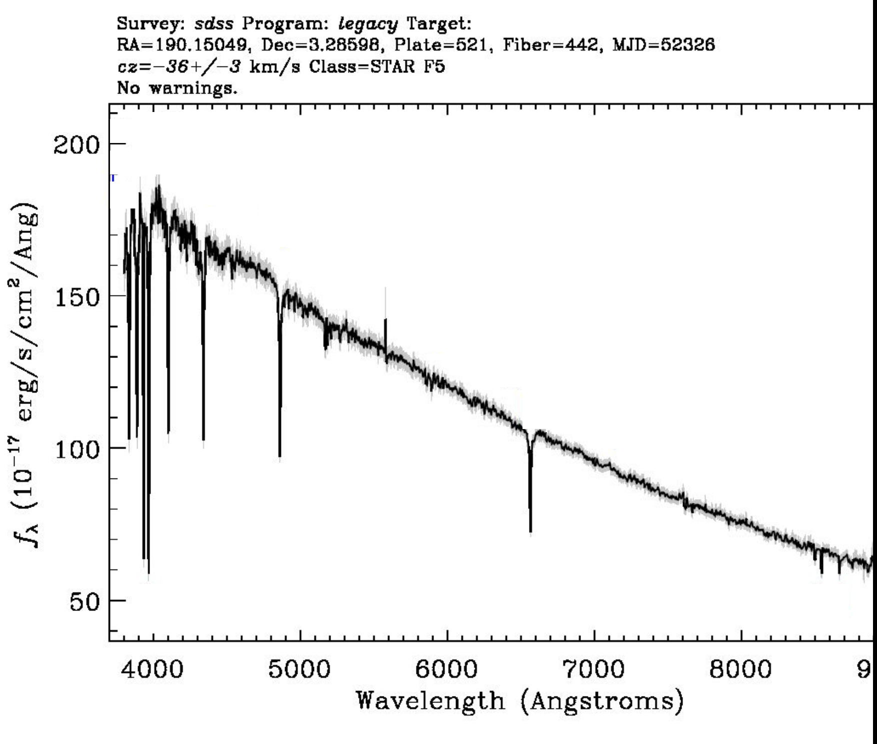
Too many stars!

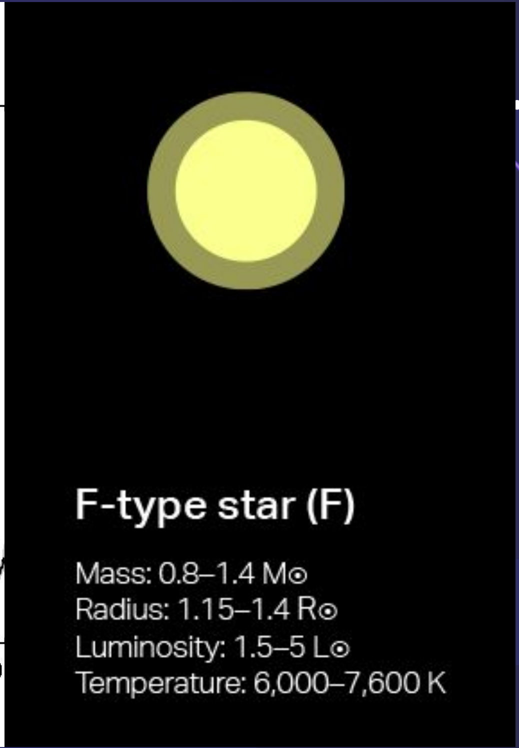
⇩

Automate the process

Sample Pipeline







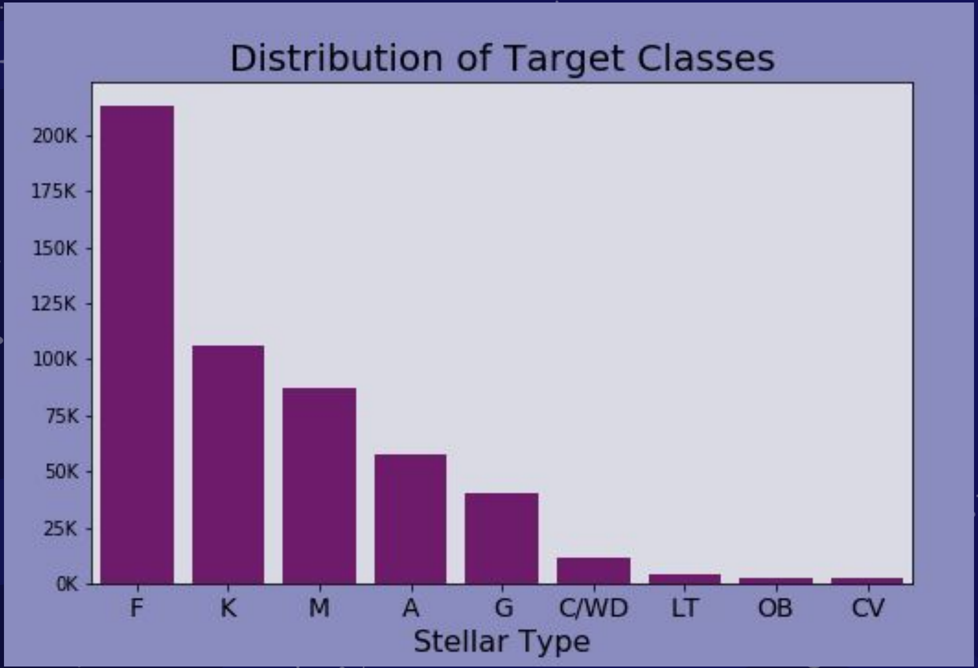
I chose this model because:

Model

Histogram-based Gradient Boosting Classification Tree

- ★ Histogram binning ➜ scales well to many features
- ★ Built in NaN handling
- ★ Much faster than other boosting models

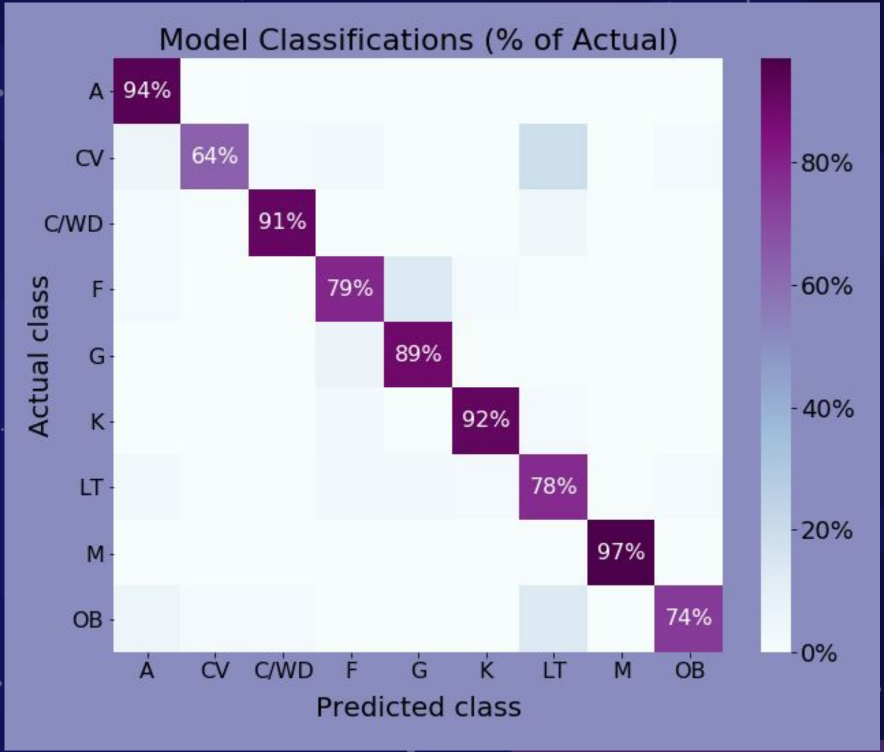
Class Distribution



- ★ Largest 5 Classes: Random Undersampling
- ★ Smallest 4 Classes: Random Oversampling

Model Performance

Balanced accuracy score: 0.85



Correct classifications are on the diagonal

Next Steps

- Expand identification capabilities
  - Subclass
  - Non-stellar objects
- Train on data from other sources

Thank You For Your Time

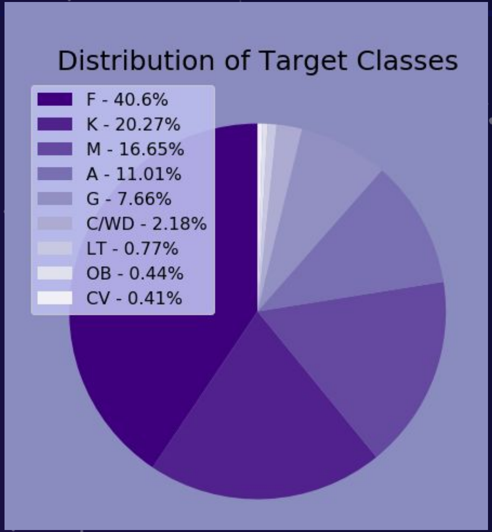
The Fine Print

This research makes use of the SciServer science platform (www.sciserver.org).

SciServer is a collaborative research environment for large-scale data-driven science. It is being developed at, and administered by, the Institute for Data Intensive Engineering and Science at Johns Hopkins University. SciServer is funded by the National Science Foundation through the Data Infrastructure Building Blocks (DIBBs) program and others, as well as by the Alfred P. Sloan Foundation and the Gordon and Betty Moore Foundation.

Appendix

Data Structure



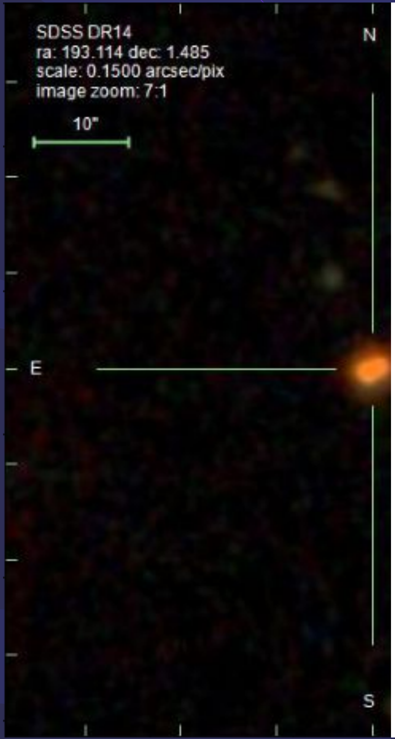
Features

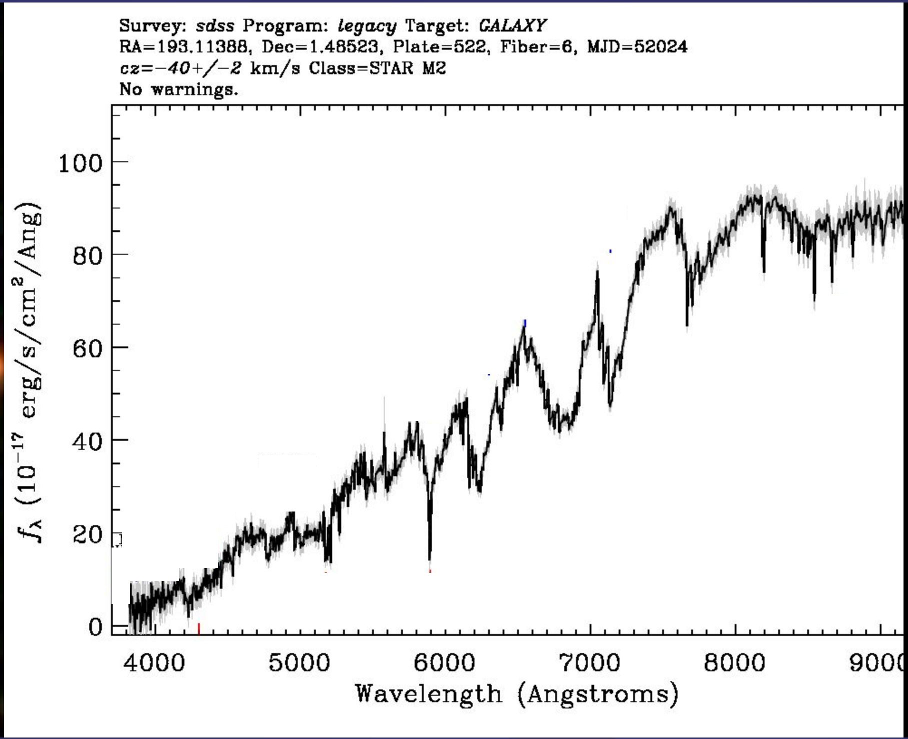
- 71 features with 4 sub features each
- Each feature represent elemental line index
- Sub features: global continuum fit, local value fit, error, pixel quality mask

Target

- Initially, 21 classes
- After grouping by subclass, only 9 classes
- Highly imbalanced: 5 of the 9 classes contain >96% of samples

Sample Pipeline 2





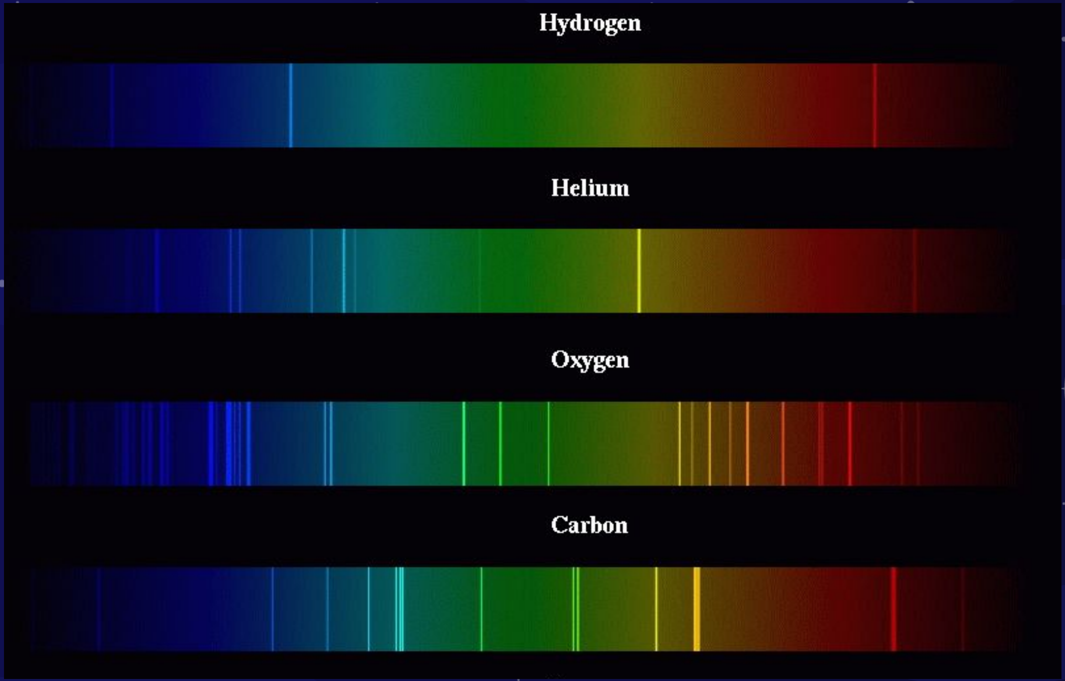
Red dwarf (M)

Mass: 0.075–0.5 M⊙
Radius: 0.08–0.45 R⊙
Luminosity: 0.015–0.08 L⊙
Temperature: 2,300–4,000 K

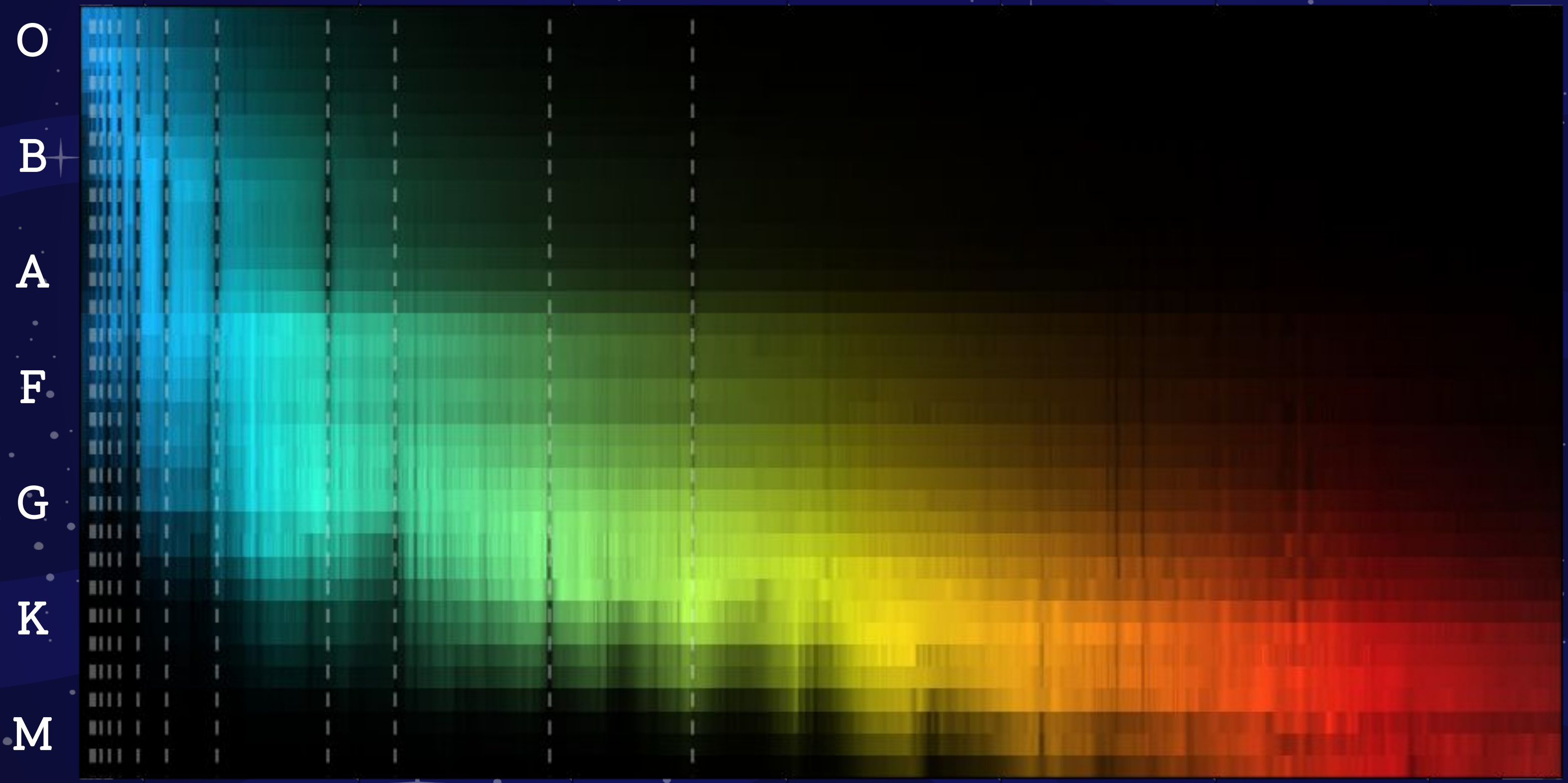
More of the Science

Spectrography

- Different elements absorb and emit different wavelengths of light
- Therefore we can determine presence of elements by looking for absorption lines (dark lines) and emission lines (bright lines)



Spectrography

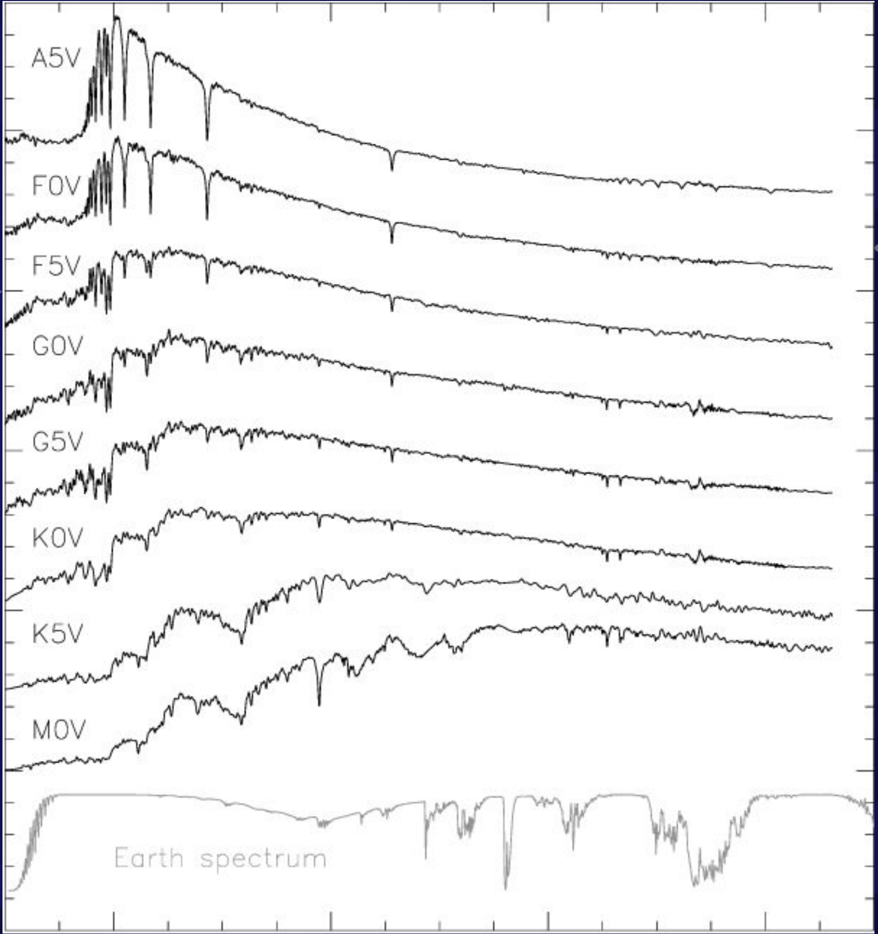


A bunch of different stellar spectra, ranging from Type O to Type M.

The different elements in these stars are causing lots of little lines.

Spectrographs

- These spectrographs can also be represented as graphs like these.
- Those absorptions and emission lines now show up as sharp spikes or dips on the graph.
- The surface temperature of the star can be determined from these graphs by looking at highest peak.



Spectral Types

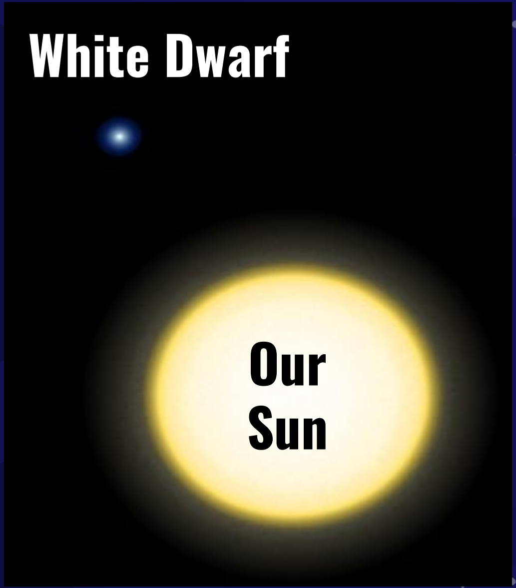
Normal spectral types range from blue hypergiants in Type O to brown dwarfs in Y (which emit almost no light).

Spectral Types

The Itsy Bitsy Ones

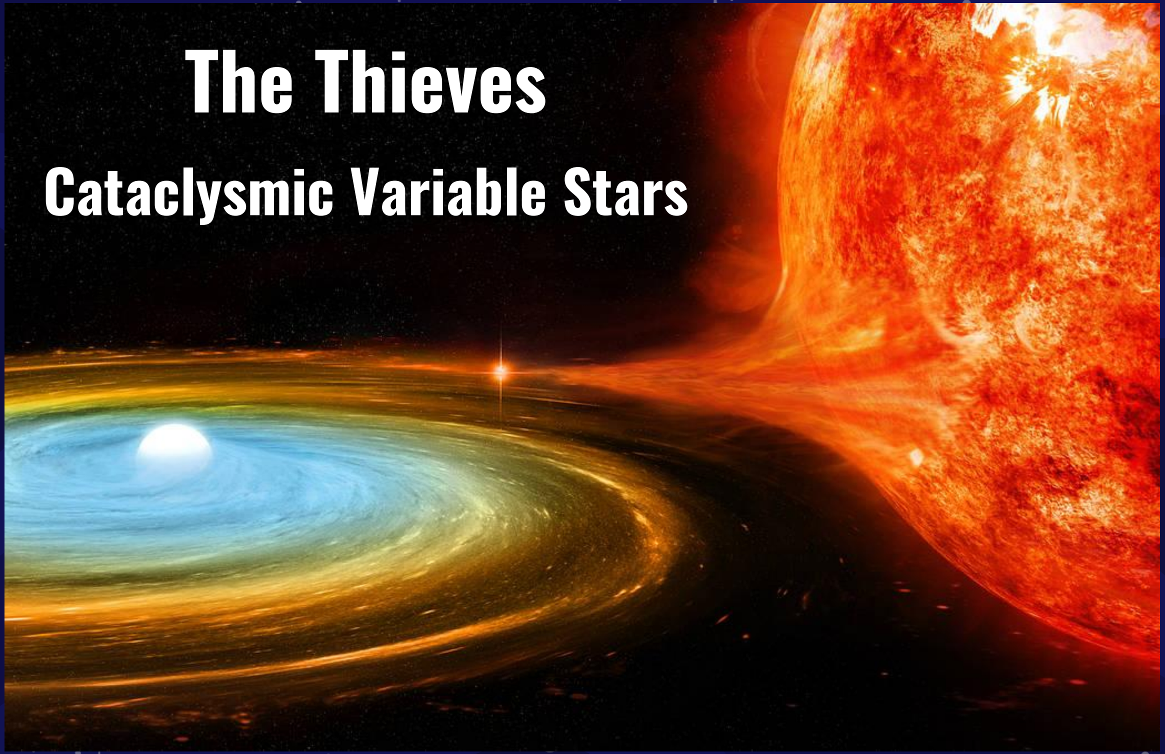
But there are also abnormal spectral types, such as white dwarfs, relatively bright but very small stars

Usually these are the core of a larger star that has blown off its outer layers in a supernova



Spectral Types

Cataclysmic variable stars are a type of star that fluctuates in brightness over time. These stars only appear in binary systems

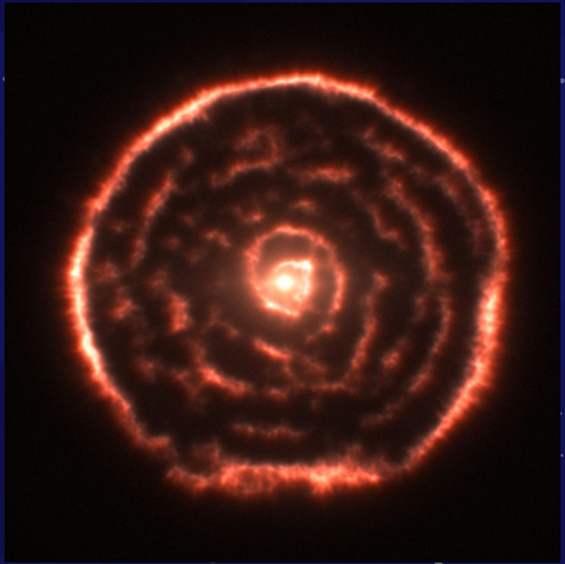


CV stars siphon material off of their binary companion, periodically building up layers and then blowing them off in a bright flash

Spectral Types

Carbon stars are the faint remnants of late stage (dying) stars that didn’t have enough mass to supernova, so once fusion stops in the core, the outer layers gradually drift off

The Dying Giants



Carbon Stars