

Classical Mechanics

Lecture notes from the second-year undergraduate course *Klasična mehanika* (Classical Mechanics) given by professor Anton Ramšak at the Faculty of Mathematics and Physics at the University of Ljubljana in the academic year 2019-20. The course covers Lagrangian and Hamiltonian mechanics at an undergraduate level. These notes do not reflect the lectures verbatim, and also draw from Professor David Tong's [lecture notes on Classical Dynamics](#) and Hebert Goldstein's *Classical Mechanics*.

Disclaimer: This document will inevitably contain some mistakes—both simple typos and legitimate errors. Keep in mind that these are the notes of an undergraduate student in the process of learning the material himself, so take what you read with a grain of salt. If you find mistakes and feel like telling me, I will be grateful and happy to hear from you, even for the most trivial of errors. You can reach me by email, in English, Slovene, or Spanish, at ejmastnak@gmail.com.

For more notes like this, visit ejmastnak.github.io/fmf

Elijan J. Mastnak
Summer Semester 2019-2020
Last update: May 31, 2022
Faculty of Mathematics and Physics, University of Ljubljana

Contents

1	Curvilinear Coordinates and Newtonian Dynamics	5
1.1	Curvilinear Coordinate Systems	5
1.1.1	Cartesian Coordinates	5
1.1.2	General Curvilinear Coordinates in Three Dimensions	5
1.1.3	Polar Coordinates	6
1.1.4	Spherical Coordinates	7
1.2	Dynamics of a Single Particle	7
1.2.1	Reference Frames	7
1.2.2	Momentum and Force	7
1.2.3	Angular Momentum and Torque	8
1.2.4	Work, Energy and Conservative Forces	8
1.3	Dynamics of Systems of Particles	9
1.3.1	Center of Mass	9
1.3.2	Momentum and Forces	9
1.3.3	Angular Momentum and Torque	9
1.3.4	Work and Energy	10
1.3.5	Virial Theorem	11
1.4	Newton's Laws in Rotating Systems	12
1.4.1	Velocity and Acceleration in a Rotating Frame	13
1.4.2	Newton's Second Law in Rotating Frames	14
1.4.3	Fictitious Forces	15
1.4.4	Rotation With Moving Origins	16
2	Lagrangian Mechanics	17
2.1	Lagrange's Equations and the Principle of the Least Action	17
2.1.1	Coordinates and Degrees of Freedom	17
2.1.2	Lagrangian and Principle of Least Action	17
2.1.3	Proof of Least Action Principle	18
2.1.4	Lagrange Equations In Changing Coordinate Systems	19
2.2	Constraints and Generalized Coordinates	21
2.2.1	Holonomic Constraints	21
2.2.2	Holonomic Constraints and Generalized Coordinates	23
2.2.3	A Summary of Constraints and Coordinates in Lagrangian Mechanics	23
2.3	Lagrange's Equations and D'Alembert's Principle	24
2.3.1	Virtual Displacements and D'Alembert's Principle	24
2.3.2	Deriving the Lagrange Equations: Take Two, Part 1	26
2.3.3	Deriving the Lagrange Equations: Take Two, Part Two	29
2.4	Noether's Theorem and Symmetries of the Lagrangian	29
2.4.1	Constants of Motion and Some Examples	29
2.4.2	Symmetries of the Lagrangian	30
2.4.3	Statement of Noether's Theorem	31
2.4.4	Examples with Noether's Theorem	32
2.5	Small Oscillations and Stability	34
2.5.1	One Dimensional Example	34
2.5.2	N Degrees of Freedom	35
2.5.3	Direct Lagrangian Mechanics Approach	38
2.5.4	The Simplified Eigenvalue Problem	40
2.5.5	Generalized Eigenvalue Problem	41

2.5.6	Normal Coordinates	42
2.6	A Concluding Example: A Particle in an Electromagnetic Field	43
3	Dynamics of Central Force Motion	44
3.1	The Central Force Problem	44
3.1.1	Two Body Problem	44
3.1.2	Basics of the Central Force Problem	45
3.1.3	Polar Coordinates and the Central Force Problem	45
3.1.4	More On Central Forces	47
3.1.5	Effective Potential	48
3.1.6	Orbits with $V(r) = -\frac{k}{r}$ Classified by energy	48
3.1.7	Existence and Stability of Circular Orbits	49
3.2	The Orbit Equation	50
3.2.1	The Kepler Problem	51
3.2.2	Analyzing the Solutions	52
3.2.3	Energy of the Orbit Revisited	53
3.2.4	Kepler's Laws of Planetary Motion	53
3.2.5	The Two-Body Central Force Problem with Lagrangian Mechanics	55
3.3	Laplace-Runge-Lenz Vector	57
3.3.1	Derivation	57
3.3.2	The LRL Vector and the Orbit Equation	58
3.4	Scattering	58
4	The Motion of Rigid Bodies	59
4.1	Rigid Body Kinematics	59
4.1.1	On Frames and Degrees of Freedom	59
4.1.2	Angular Velocity	59
4.1.3	Inertia Tensor	61
4.1.4	Parallel Axis Theorem	62
4.1.5	Angular Momentum	62
4.1.6	Rigid Body Motion as Translation and Rotation	63
4.1.7	Euler Equations	63
4.2	Motion of the Free Top	65
4.2.1	Symmetric Free Top	65
4.2.2	Stability and Perturbations of The Asymmetric Top	66
4.3	Euler Angles	66
4.3.1	Angular Velocity and Kinetic Energy	67
4.3.2	Free Symmetric Top In Terms of Euler Angles	68
4.4	Heavy Symmetric Top	68
4.4.1	Lagrangian and Conserved Quantities	69
4.4.2	Letting the Top Go	70
4.4.3	Uniform Precession Without Nutation	70
4.4.4	The Sleeping Top	71
5	Hamiltonian Mechanics	73
5.1	Hamilton's Equations	73
5.1.1	Mathematical Detour: The Legendre Transform	73
5.1.2	Hamilton's Equations of Motion	74
5.1.3	Examples Using the Hamiltonian Equations of Motion	75
5.1.4	Hamiltonian Mechanics and Conservation Laws	77

5.1.5	Principle of Least Action á la Hamilton	77
5.2	Poisson Brackets	78
5.2.1	Definition and Properties	78
5.2.2	Poisson Brackets and Angular Momentum	80

1 Curvilinear Coordinates and Newtonian Dynamics

1.1 Curvilinear Coordinate Systems

In general, we are interested in:

1. Coordinates and coordinate curves
2. Expression for position \mathbf{r} in terms of coordinates
3. Expressions for basis vectors
4. Magnitudes of each basis vector $|\mathbf{e}_i| = \sqrt{|\mathbf{e}_i \cdot \mathbf{e}_i|}$ as well as dot products in mixed basis vector terms. We are basically checking if the basis vectors are orthonormal.

1.1.1 Cartesian Coordinates

- Coordinates: (x, y, z)
- Basis vectors: $\hat{\mathbf{i}}, \hat{\mathbf{j}}, \hat{\mathbf{k}}$
- Position: $\mathbf{r} = x\hat{\mathbf{i}} + y\hat{\mathbf{j}} + z\hat{\mathbf{k}}$
- Definitions of basis vectors:

$$\hat{\mathbf{i}} = \frac{\partial \mathbf{r}}{\partial x} \qquad \hat{\mathbf{j}} = \frac{\partial \mathbf{r}}{\partial y} \qquad \hat{\mathbf{k}} = \frac{\partial \mathbf{r}}{\partial z}$$

In Cartesian coordinates, the basis vectors have the same magnitude and direction at all points in space. For this reason, we call $\hat{\mathbf{i}}, \hat{\mathbf{j}}, \hat{\mathbf{k}}$ a *global basis*.

- The global Cartesian basis $\hat{\mathbf{i}}, \hat{\mathbf{j}}, \hat{\mathbf{k}}$ is orthonormal. This means that all the basis vectors have magnitude one while the dot products of mixed basis are zero.

$$\begin{aligned} |\hat{\mathbf{i}}| &= \sqrt{\hat{\mathbf{i}} \cdot \hat{\mathbf{i}}} = 1 & |\hat{\mathbf{j}}| &= \sqrt{\hat{\mathbf{j}} \cdot \hat{\mathbf{j}}} = 1 & |\hat{\mathbf{k}}| &= \sqrt{\hat{\mathbf{k}} \cdot \hat{\mathbf{k}}} = 1 \\ \hat{\mathbf{i}} \cdot \hat{\mathbf{j}} &= \hat{\mathbf{j}} \cdot \hat{\mathbf{i}} = 0 & \hat{\mathbf{i}} \cdot \hat{\mathbf{k}} &= \hat{\mathbf{k}} \cdot \hat{\mathbf{i}} = 0 & \hat{\mathbf{j}} \cdot \hat{\mathbf{k}} &= \hat{\mathbf{k}} \cdot \hat{\mathbf{j}} = 0 \end{aligned}$$

1.1.2 General Curvilinear Coordinates in Three Dimensions

For a general three-dimensional coordinate system:

1. Coordinates: $(x_\alpha, x_\beta, x_\gamma)$
2. Basis vectors: $\mathbf{e}_\alpha, \mathbf{e}_\beta, \mathbf{e}_\gamma$
3. Definitions of basis vectors:

$$\mathbf{e}_\alpha = \frac{\partial \mathbf{r}}{\partial x_\alpha} \qquad \mathbf{e}_\beta = \frac{\partial \mathbf{r}}{\partial x_\beta} \qquad \mathbf{e}_\gamma = \frac{\partial \mathbf{r}}{\partial x_\gamma}$$

Important: In general, the basis vectors $\mathbf{e}_\alpha, \mathbf{e}_\beta, \mathbf{e}_\gamma$ vary in magnitude and direction at different points in space. Because $\mathbf{e}_\alpha, \mathbf{e}_\beta, \mathbf{e}_\gamma$ are different at different points in space, we cannot define position \mathbf{r} in terms of $\mathbf{e}_\alpha, \mathbf{e}_\beta, \mathbf{e}_\gamma$. This wouldn't be well defined. Instead, we express position in terms of the global basis $\hat{\mathbf{i}}, \hat{\mathbf{j}}, \hat{\mathbf{k}}$ and write the Cartesian coordinates x, y, z as functions of the curvilinear coordinates $x_\alpha, x_\beta, x_\gamma$.

4. Position: $\mathbf{r} = x(x_\alpha, x_\beta, x_\gamma) \hat{\mathbf{i}} + y(x_\alpha, x_\beta, x_\gamma) \hat{\mathbf{j}} + z(x_\alpha, x_\beta, x_\gamma) \hat{\mathbf{k}}$

We are interested in how the Cartesian coordinates (x, y, z) are expressed in terms of the curvilinear coordinates $(x_\alpha, x_\beta, x_\gamma)$. That is, we want expressions for x, y , and z as functions of $(x_\alpha, x_\beta, x_\gamma)$.

5. Differential of position and velocity:

$$\begin{aligned} d\mathbf{r} &= \mathbf{e}_\alpha dx_\alpha + \mathbf{e}_\beta dx_\beta + \mathbf{e}_\gamma dx_\gamma \\ \mathbf{v} &= \frac{d\mathbf{r}}{dt} = \mathbf{e}_\alpha \frac{dx_\alpha}{dt} + \mathbf{e}_\beta \frac{dx_\beta}{dt} + \mathbf{e}_\gamma \frac{dx_\gamma}{dt} \\ &= \mathbf{e}_\alpha \dot{x}_\alpha + \mathbf{e}_\beta \dot{x}_\beta + \mathbf{e}_\gamma \dot{x}_\gamma \end{aligned}$$

Locally, the general basis vectors $\mathbf{e}_\alpha, \mathbf{e}_\beta, \mathbf{e}_\gamma$ are constant in a small neighborhood of every point in space. Because of this local constant behavior on a small scale, we can write the differential of position in terms of general basis vectors.

1.1.3 Polar Coordinates

Polar coordinates are used in two-dimensional space in situations with circular symmetry.

1. Coordinates: (r, ϕ)

- (a) r is the distance from a point in space to the origin
- (b) ϕ is the angle between the basis vector $\hat{\mathbf{i}}$ and a ray connected a point to the origin.
- (c) The coordinate curves for r are concentric circles centered at the origin, while the coordinate curves for ϕ are rays originating at the origin radiating out into space.

2. Position:

$$\mathbf{r}(r, \phi) = r \cos \phi \hat{\mathbf{i}} + r \sin \phi \hat{\mathbf{j}} \quad \text{where} \quad x(r, \phi) = r \cos \phi \quad y(r, \phi) = r \sin \phi$$

The Jacobian determinant for polar coordinates is $\det \mathbf{J} = r$.

3. Basis vectors:

$$\mathbf{e}_r = \cos \phi \hat{\mathbf{i}} + \sin \phi \hat{\mathbf{j}} \quad \mathbf{e}_\phi = -r \sin \phi \hat{\mathbf{i}} + r \cos \phi \hat{\mathbf{j}}$$

4. The polar basis is orthogonal but not fully normal. We have:

$$|\mathbf{e}_r|^2 = 1 \quad |\mathbf{e}_\phi|^2 = r^2 \quad \mathbf{e}_r \cdot \mathbf{e}_\phi = 0$$

5. In polar coordinates, kinetic energy T , acceleration $\ddot{\mathbf{r}}$, and Newton's second law are:

$$T = \frac{1}{2} m (\dot{r}^2 + r^2 \dot{\phi}^2) \quad \sum \mathbf{F} = m \ddot{\mathbf{r}} = m \mathbf{e}_r (\ddot{r} + \dot{\phi}^2 r) + m \mathbf{e}_\phi \left(\ddot{\phi} + 2 \frac{\dot{\phi} \dot{r}}{r} \right)$$

1.1.4 Spherical Coordinates

1. Coordinates: (r, ϕ, θ) where r is the distance from a point in space to the origin, ϕ is the angle between $\hat{\mathbf{i}}$ and the projection of the point onto the $\hat{\mathbf{i}}\hat{\mathbf{j}}$ plane, and θ is the angle between the point and $\hat{\mathbf{k}}$.

The coordinate curves for r are concentric spheres centered at the origin, while the coordinate curves for ϕ and θ are rays radiating from the origin into space.

2. Position:

$$\begin{aligned}\mathbf{r}(r, \phi) &= r \cos \phi \sin \theta \hat{\mathbf{i}} + r \sin \phi \sin \theta \hat{\mathbf{j}} + r \cos \theta \hat{\mathbf{k}} \\ x(r, \phi, \theta) &= r \cos \phi \sin \theta \quad y(r, \phi, \theta) = r \sin \phi \sin \theta \quad z(r, \phi, \theta) = r \cos \theta\end{aligned}$$

The Jacobian determinant for spherical coordinates is $\det \mathbf{J} = r^2 \sin \theta$.

3. Basis vectors:

$$\begin{aligned}\mathbf{e}_r &= \cos \phi \sin \theta \hat{\mathbf{i}} + \sin \phi \sin \theta \hat{\mathbf{j}} + \cos \theta \hat{\mathbf{k}} \\ \mathbf{e}_\phi &= -r \sin \phi \sin \theta \hat{\mathbf{i}} + r \cos \phi \sin \theta \hat{\mathbf{j}} + 0 \hat{\mathbf{k}} \\ \mathbf{e}_\theta &= -r \cos \phi \cos \theta \hat{\mathbf{i}} + r \sin \phi \cos \theta \hat{\mathbf{j}} - r \sin \theta \hat{\mathbf{k}}\end{aligned}$$

4. The spherical basis is orthogonal but not fully normal. We have:

$$|\mathbf{e}_r|^2 = 1 \quad |\mathbf{e}_\phi|^2 = r^2 \sin^2 \theta \quad |\mathbf{e}_\theta|^2 = r^2$$

5. In spherical coordinates, kinetic energy T is

$$T = \frac{1}{2}m \left(\dot{r}^2 + \dot{\phi}^2 r^2 \sin^2 \theta + \dot{\theta}^2 r^2 \right)$$

1.2 Dynamics of a Single Particle

A particle is an object of insignificant size, relative to the context of the problem.

1.2.1 Reference Frames

- An inertial frame is unaccelerated.

Equivalently, an inertial frame is defined as a frame in which a free particle with constant mass travels in straight line $\mathbf{r} = \mathbf{r}_0 + \mathbf{v}t$.

- In Euclidean space, there are 10 linearly independent transformations that preserve inertial frames. These are three rotations, three translations, three boosts and one time translation, and together they form the Galilean group.

1.2.2 Momentum and Force

- A particle's *linear momentum* \mathbf{p} is defined as $\mathbf{p} = m\mathbf{v}$.
- Linear momentum and force are independent of the choice of origin.
- The general form of Newton's second law reads $\mathbf{F}(\mathbf{r}, \dot{\mathbf{r}}) = \frac{d\mathbf{p}}{dt}$. For a particle with constant mass, $\mathbf{F} = m\ddot{\mathbf{r}}$

- If the net force \mathbf{F} on a particle is zero, then the particle's linear momentum \mathbf{p} is conserved. $\mathbf{F} = 0 \implies \mathbf{p} = \text{constant}$.
- Newton's third law states $\mathbf{F}_{12} = -\mathbf{F}_{21}$

Newton's laws only hold in inertial frames.

1.2.3 Angular Momentum and Torque

- A particle's *angular momentum* \mathbf{L} with respect to the origin is defined as $\mathbf{L} = \mathbf{r} \times \mathbf{p} = m(\mathbf{r} \times \mathbf{v})$
- The rotational analog of Newton's second law relates the torque $\boldsymbol{\tau}$ on the particle relative to the origin to angular momentum \mathbf{L} with respect to the origin via $\boldsymbol{\tau} = \frac{d\mathbf{L}}{dt}$.

Torque is related to force by $\boldsymbol{\tau} = \mathbf{r} \times \mathbf{F}$, obtained via

$$\boldsymbol{\tau} = \frac{d\mathbf{L}}{dt} = \frac{d}{dt}[\mathbf{r} \times \mathbf{p}] = \dot{\mathbf{r}} \times \mathbf{p} + \mathbf{r} \times \dot{\mathbf{p}} = \mathbf{0} + \mathbf{r} \times \mathbf{F} = \mathbf{r} \times \mathbf{F}$$

since $\dot{\mathbf{r}} \times \mathbf{p} = m(\mathbf{v} \times \mathbf{v}) = m \cdot (\mathbf{0}) = \mathbf{0}$.

- If the net torque $\boldsymbol{\tau}$ on a particle is zero, then the particle's angular momentum \mathbf{L} is conserved. $\boldsymbol{\tau} = 0 \implies \mathbf{L} = \text{constant}$.
- Both angular momentum and torque are measured with respect to a chosen reference point and thus depend on the choice of origin.

1.2.4 Work, Energy and Conservative Forces

- The work W done by a force \mathbf{F} on a particle along the curve γ is defined by the line integral $W = \int_{\gamma} \mathbf{F} \cdot d\mathbf{r}$.
- The *kinetic energy* of a particle of mass m moving with velocity \mathbf{v} is defined as $T = \frac{1}{2}m|\mathbf{v}|^2 = \frac{1}{2}mv^2$.
- The change in a particle's kinetic energy under the influence of a force \mathbf{F} equals the work done by the force.

If the particle travels from position \mathbf{r}_1 at time t_1 to position \mathbf{r}_2 at time t_2 , the change in the particle's kinetic energy is

$$T(t_2) - T(t_1) = \int_{t_1}^{t_2} \frac{dT}{dt} dt = \int_{t_1}^{t_2} \mathbf{F} \cdot \dot{\mathbf{r}} dt = \int_{\mathbf{r}_1}^{\mathbf{r}_2} \mathbf{F} \cdot d\mathbf{r} = W$$

- *Conservative forces* are a special type of force that depend only on \mathbf{r} and not $\dot{\mathbf{r}}$ and have the property that the work done by the force on a particle moving between two points is independent of the path taken.

Conservative forces can be written as the gradient of a potential energy V . In this case, $\mathbf{F} = -\nabla V$. The work done by a conservative force on a particle along a curve γ equals the change in potential energy between the endpoints.

$$W = \int_{\gamma} \mathbf{F} \cdot d\mathbf{r} = -(V_2 - V_1) \quad \text{and} \quad \oint_{\kappa} \mathbf{F} \cdot d\mathbf{r} = 0$$

where κ is a closed curve.

- Conservative forces arising from potentials include the gravitational and electrostatic forces. For particles under the influence of conservative forces, we have the energy conservation law

$$T_1 + V_1 = T_2 + V_2 \equiv E$$

stating that the particle's total energy E is conserved. The law arises from

$$W = T_2 - T_1 = - \int_{\mathbf{r}_1}^{\mathbf{r}_2} V(r) \cdot d\mathbf{r} = -V_2 + V_1$$

1.3 Dynamics of Systems of Particles

This section considers a system of N particles with masses m_1, m_2, \dots, m_N and position vectors $\mathbf{r}_1, \mathbf{r}_2, \dots, \mathbf{r}_N$.

1.3.1 Center of Mass

- The total mass of the system is $M = m_1 + m_2 + \dots + m_N$
- The system's center of mass \mathbf{r}^* is

$$\mathbf{r}^* = \frac{m_1 \mathbf{r}_1 + m_2 \mathbf{r}_2 + \dots + m_N \mathbf{r}_N}{m_1 + m_2 + \dots + m_N} = \frac{\sum_{i=1}^N m_i \mathbf{r}_i}{\sum_{i=1}^N m_i} = \frac{1}{M} \sum_{i=1}^N m_i \mathbf{r}_i$$

- The velocity of the center of mass is $\mathbf{v}^* = \frac{d\mathbf{r}^*}{dt}$, and the acceleration of the center of mass is $\mathbf{a}^* = \frac{d\mathbf{v}^*}{dt} = \frac{d^2 \mathbf{r}^*}{dt^2}$

1.3.2 Momentum and Forces

- The momentum \mathbf{p}^* of the system's center of mass is

$$\mathbf{p}^* = \sum_{i=1}^N \mathbf{p}_i = M \frac{d\mathbf{r}^*}{dt} = M \mathbf{v}^*$$

- The net external force \mathbf{F}_{ext} acting on the system is related to the center of mass's momentum via $\mathbf{F}_{ext} = \frac{d\mathbf{p}^*}{dt}$.

If the system's mass is constant, $\mathbf{F}_{ext} = \frac{d\mathbf{p}^*}{dt} = M \mathbf{a}^*$

- Central internal forces are forces that can be written in the form

$$\mathbf{F}_{ij} = F_{ij} \frac{\mathbf{r}_i - \mathbf{r}_j}{|\mathbf{r}_i - \mathbf{r}_j|} \quad i \neq j$$

1.3.3 Angular Momentum and Torque

- The angular momentum \mathbf{L}^* of the system's center of mass relative to the origin is

$$\mathbf{L}^* = \sum_{i=1}^N \mathbf{L}_i = M(\mathbf{r}^* \times \mathbf{v}^*) + \sum_{i=1}^N m_i(\mathbf{r}'_i \times \mathbf{v}'_i)$$

where \mathbf{r}'_i is the position of the i th particle relative to the center of mass \mathbf{r}^* , defined implicitly by the relationship $\mathbf{r}_i = \mathbf{r}^* + \mathbf{r}'_i$.

- The angular momentum \mathbf{L}^* of the center of mass is related to the net torque $\boldsymbol{\tau}$ acting on the system via

$$\boldsymbol{\tau}_{ext} + \boldsymbol{\tau}_{int} = \frac{d\mathbf{L}^*}{dt}$$

If all internal forces are central, the net internal torque vanishes and we have

$$\boldsymbol{\tau}_{ext} = \frac{d\mathbf{L}^*}{dt}$$

1.3.4 Work and Energy

- The total kinetic energy of a system of N particles is

$$T = \frac{1}{2} \sum_i^N m_i \dot{\mathbf{r}}_i^2$$

We can decompose the position vector \mathbf{r}_i into

$$\mathbf{r}_i = \mathbf{R} + \tilde{\mathbf{r}}_i$$

where \mathbf{R} is the position of the system's center of mass and $\tilde{\mathbf{r}}_i$ is the distance from the i th particle to the center of mass.

- With the decomposition $\mathbf{r}_i = \mathbf{R} + \tilde{\mathbf{r}}_i$, the total kinetic energy becomes

$$T = \frac{1}{2} M \dot{\mathbf{R}}^2 + \frac{1}{2} \sum_i^N m_i \dot{\tilde{\mathbf{r}}}_i^2$$

This form tells us that the total kinetic interpretation splits up into the kinetic energy of the center of mass and the *internal kinetic energy* describing how the system's particles move around the center of mass.

- Following the same story line as for a single particle, the change in kinetic energy between two times is

$$T(t_2) - T(t_1) = \sum_i^N \int \mathbf{F}_i^{\text{ext}} \cdot d\mathbf{r}_i + \sum_{i \neq j} \int \mathbf{F}_{ij} \cdot d\mathbf{r}_i$$

As before, we need conservative forces to recover energy conservation. For systems of particles, we need both

- Conservative external forces $\mathbf{F}_i^{\text{ext}} = -\nabla_i V_i(\mathbf{r}_1, \dots, \mathbf{r}_N)$ and...
- Conservative internal forces $\mathbf{F}_{ij} = -\nabla_i V_{ij}(\mathbf{r}_1, \dots, \mathbf{r}_N)$

where $\nabla_i \equiv \frac{\partial}{\partial \mathbf{r}_i}$.

- To recover Newton's third law $\mathbf{F}_{ij} = -\mathbf{F}_{ji}$, along with the condition that \mathbf{F}_{ij} is parallel to the relative position $\mathbf{r}_i - \mathbf{r}_j$, we must consider central internal potentials satisfying

$$V_{ij} = V_{ji} \quad \text{and} \quad V_{ij} = V_{ij}(|\mathbf{r}_i - \mathbf{r}_j|)$$

More so, must consider a restriction on the external forces of the form

$$V_i = V_i(\mathbf{r}_i)$$

so the force on the i th particle is independent of the positions of the other particles.

With these conditions, we can define the *total potential energy*

$$V = \sum_i^N V_i + \sum_{i < j} V_{ij}$$

and show that the system's *total energy* $H = T + V$ is conserved. **TODO**

1.3.5 Virial Theorem

- We consider a stable system of N particles and define the quantity

$$G = \sum_i^N \mathbf{p}_i \cdot \mathbf{r}_i$$

Using the product rule, the rate of change of G with respect to time is

$$\frac{dG}{dt} = \sum_i \dot{\mathbf{p}}_i \cdot \mathbf{r}_i + \sum_i \mathbf{p}_i \cdot \dot{\mathbf{r}}_i = \sum_i \mathbf{F}_i \cdot \mathbf{r}_i + 2T$$

where $T = \frac{1}{2} \sum_i m \dot{\mathbf{r}}^2$ is the system's total kinetic energy.

- Here's the key step: as long as the system is stable, it will stay in a bounded region of phase space, so G can only take on a bounded set of values. Because the value of G is bounded, the average value of its rate of change $\frac{dG}{dt}$ over an infinite time period is zero.

If we can further assume the system makes periodic trajectory in phase space, we can reach the same conclusion more elegantly. Suppose the system starts at some point in its phase space where $G = G_0$. Because of the imposed periodicity, after some amount of time the system will eventually return to its starting point in phase space and G will return to its initial value G_0 . The average of $\frac{dG}{dt}$ over this time period vanishes.

In either case, we have the result $\langle \frac{dG}{dt} \rangle = 0$ over a suitably long time period.

- Taking the time average of the equation for $\frac{dG}{dt}$ and applying $\langle \frac{dG}{dt} \rangle = 0$ gives

$$\left\langle \frac{dG}{dt} \right\rangle = \left\langle \sum_i \mathbf{F}_i \cdot \mathbf{r}_i \right\rangle + 2 \langle T \rangle = 0 \implies 2 \langle T \rangle = - \left\langle \sum_i \mathbf{F}_i \cdot \mathbf{r}_i \right\rangle$$

This is the *virial theorem* for a stable system of N particles. What exactly does it mean? In fact, it is a way of relating the system's average kinetic and potential energies. The potential energy is hidden in the \mathbf{F}_i : if we can assume the forces are conservative, we can relate force to potential with $\mathbf{F}(\mathbf{r}) = -\nabla V(\mathbf{r})$.

- The virial theorem gives a beautiful result for interactions governed by a central potential $V = V(r)$ of the form $V(r) = kr^n$; the result is

$$2 \langle T \rangle = n \langle V \rangle$$

To derive this, we start with an expression for the force. Although the discussion is a bit premature (see the central forces chapter), because the force is central we can write

$$\mathbf{F} = -\nabla V(r) = -\nabla(kr^n) = -knr^{n-1} \hat{\mathbf{r}}$$

where $\hat{\mathbf{r}} = \frac{\mathbf{r}}{r}$. In component form, the force on the i th particle is $\mathbf{F}_i = -knr_i^{n-1} \hat{\mathbf{r}}_i$.

$$\mathbf{F}_i = -knr_i^{n-1} \hat{\mathbf{r}}_i = -knr_i^{n-2} \mathbf{r}_i \quad \text{where} \quad \hat{\mathbf{r}}_i = \frac{\mathbf{r}_i}{r_i}$$

Evaluating the dot product $\mathbf{F}_i \cdot \mathbf{r}_i$ in the sum gives

$$\mathbf{F}_i \cdot \mathbf{r}_i = (-knr_i^{n-2} \mathbf{r}_i) \cdot \mathbf{r}_i = -knr_i^n = -nV(r_i)$$

All we have to do is sum over the $i = 1, \dots, N$ particles to get

$$-\left\langle \sum_i \mathbf{F}_i \cdot \mathbf{r}_i \right\rangle = -\left\langle \sum_i (-nV(r_i)) \right\rangle = +nV(r) \implies 2\langle T \rangle = n\langle V \rangle$$

1.4 Newton's Laws in Rotating Systems

This section will generalize Newton's second law $\mathbf{F} = \dot{\mathbf{p}}$, which is valid only in inertial frames, to a more general expression that holds in both inertial and non-inertial frames. If you think about it, this an important goal, since nearly every frame of reference you will work with is to some extent non-inertial (the Earth, for instance rotates around both its own axes and the sun).

- In this section we will be dealing with
 - An inertial coordinate system S with coordinate axes (x, y, z)
 - An rotating coordinate system S' with coordinate axes (x', y', z')
 - A vector quantity in space (typically the position vector \mathbf{r}), which we will describe in terms of both the inertial or non-inertial coordinates.

In this first section, the origins of S and S' will coincide. S is at rest and S' rotates with respect to S about an axis through the mutual origin with angular velocity $\boldsymbol{\omega}$.

To begin, we'll assume the z and z' axes coincide and denote the angle between the x and x' axes as θ .

- We first consider the motion of particles. Consider a particle that is sitting stationary in the rotating frame S' . This could be, for example, a child sitting on a carousel. From the perspective of the fixed frame S , the particle will appear to move with the velocity

$$\dot{\mathbf{r}} = \boldsymbol{\omega} \times \mathbf{r}$$

Because we choose the z and z' axes to coincide we have $\boldsymbol{\omega} = \omega \hat{\mathbf{z}} = \dot{\theta} \hat{\mathbf{z}}$ where $\omega = \dot{\theta}$ is the angular speed and the direction of $\boldsymbol{\omega}$ is the axis of rotation, defined in the right-hand sense.

- The key step is extending this description of rotation to the axes of S' themselves. Let $\mathbf{e}'_1, \mathbf{e}'_2$ and \mathbf{e}'_3 be the unit vectors pointing along the x', y' and z' axes of S' . Just like a particle in S' appeared in S to move with velocity $\dot{\mathbf{r}} = \boldsymbol{\omega} \times \mathbf{r}$, the unit vectors rotate with velocity

$$\dot{\mathbf{e}}'_i = \boldsymbol{\omega} \times \mathbf{e}'_i$$

from the perspective of the inertial frame S . The rotation of the unit vectors encoded in the above equation will allow us to understand motion in rotating frames.

1.4.1 Velocity and Acceleration in a Rotating Frame

- Consider a particle moving around in space; we will then analyze its motion in both the S and S' frames. When I was first learning, it helped me to remember we are analyzing the *same* particle in both frames, we are just encoding its dynamics in two different coordinate systems. Or, more formally (in linear algebra terms), we are writing the same vector quantity in different bases.
- Let $\mathbf{e}_1, \mathbf{e}_2$ and \mathbf{e}_3 be the unit vectors pointing along the x, y and z axes of S . In this basis the position of the particle in S is

$$\mathbf{r} = \sum_{i=1}^3 r_i \mathbf{e}_i \equiv r_i \mathbf{e}_i$$

where the second equality adopts the shorthand summation convention. Alternatively, the position measured in S' is

$$\mathbf{r} = r'_i \mathbf{e}'_i$$

Note that the position vector \mathbf{r} is the same in both frames since the origins coincide, but the coordinates r_i and r'_i will in general differ, since they're measured with respect to different axes.

- Equipped with the position vector \mathbf{r} , we can now compute the velocity. The product rule for derivatives is essential here. Measured in frame S , the velocity is

$$\dot{\mathbf{r}} = \dot{r}_i \mathbf{e}_i + r_i \dot{\mathbf{e}}_i = \dot{r}_i \mathbf{e}_i$$

We have the simple expression $\dot{\mathbf{r}} = \dot{r}_i \mathbf{e}_i$ because the axes \mathbf{e}_i do not change with time (i.e. $\dot{\mathbf{e}}_i = 0$) since S is inertial. This won't hold in S' , where the velocity is

$$\begin{aligned} \dot{\mathbf{r}} &= \dot{r}'_i \mathbf{e}'_i + r'_i \dot{\mathbf{e}}'_i = \dot{r}'_i \mathbf{e}'_i + r'_i (\boldsymbol{\omega} \times \mathbf{e}'_i) = \dot{r}'_i \mathbf{e}'_i + \boldsymbol{\omega} \times (r'_i \mathbf{e}'_i) \\ &= \dot{r}'_i \mathbf{e}'_i + \boldsymbol{\omega} \times \mathbf{r} \end{aligned}$$

where we have used the handy expression $\dot{\mathbf{e}}'_i = \boldsymbol{\omega} \times \mathbf{e}'_i$.

- We now introduce a new notation to highlight the physics and hopefully clarify what is going on in the transitions between frames. We write the velocity of the particle seen by an observer in S as

$$\left(\frac{d\mathbf{r}}{dt} \right)_S = \dot{r}_i \mathbf{e}_i$$

Similarly, we write the velocity as seen by an observer rotating with the frame S' as

$$\left(\frac{d\mathbf{r}}{dt} \right)_{S'} = \dot{r}'_i \mathbf{e}'_i$$

- Equipped with this notation and the equation $\dot{\mathbf{r}} = \dot{r}'_i \mathbf{e}'_i + \boldsymbol{\omega} \times \mathbf{r}$, we see that observers in the two frames measure two different velocities related by

$$\left(\frac{d\mathbf{r}}{dt} \right)_S = \left(\frac{d\mathbf{r}}{dt} \right)_{S'} + \boldsymbol{\omega} \times \mathbf{r}$$

We now see the discrepancy is just the relative velocity $\mathbf{v}_{\text{rel}} = \boldsymbol{\omega} \times \mathbf{r}$ between the two frames.

- On to the particle's acceleration; we just have to continue our time differentiation of \mathbf{r} . In S the acceleration is simply

$$\ddot{\mathbf{r}} = \ddot{r}_i \mathbf{e}_i$$

where we are again helped by $\dot{\mathbf{e}}_i = 0$. In the frame S' the expression is more complicated. From the expression $\dot{\mathbf{r}} = \dot{r}'_i \mathbf{e}'_i + \boldsymbol{\omega} \times \mathbf{r}$ with $\dot{\mathbf{e}}'_i = \boldsymbol{\omega} \times \mathbf{e}'_i$ and $\mathbf{r} = r'_i \mathbf{e}'_i$, time differentiation and a few steps of algebra lead to

$$\ddot{\mathbf{r}} = \ddot{r}'_i \mathbf{e}'_i + 2\dot{r}'_i \boldsymbol{\omega} \times \mathbf{e}'_i + r'_i \dot{\boldsymbol{\omega}} \times \mathbf{e}'_i + r'_i \boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{e}'_i)$$

- We adopt a analogous notation for acceleration as we did for velocities: the acceleration of the particle seen by an observer in S is

$$\left(\frac{d^2 \mathbf{r}}{dt^2} \right)_S = \ddot{r}_i \mathbf{e}_i$$

and the acceleration as seen by an observer rotating with the frame S' is

$$\left(\frac{d^2 \mathbf{r}}{dt^2} \right)_{S'} = \ddot{r}'_i \mathbf{e}'_i$$

If we equate the expressions for $\ddot{\mathbf{r}}$ in the inertial and rotating bases and adopt the new notation, leaving the quantities \mathbf{r} and $\boldsymbol{\omega}$ in vector form (i.e. not expanding them in either the \mathbf{e}_i or \mathbf{e}'_i bases) we get the vector equation

$$\left(\frac{d^2 \mathbf{r}}{dt^2} \right)_S = \left(\frac{d^2 \mathbf{r}}{dt^2} \right)_{S'} + 2\boldsymbol{\omega} \times \left(\frac{d\mathbf{r}}{dt} \right)_{S'} + \dot{\boldsymbol{\omega}} \times \mathbf{r} + \boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r})$$

This result is important: it lets us generalize Newton's second law to rotating frames.

1.4.2 Newton's Second Law in Rotating Frames

- The messy mathematics is behind us. Let's take a look at how an observer rotating with the frame S' would see Newton's second law of motion applied to a particle moving space. In the inertial frame we have good old Newton's second law $m\ddot{\mathbf{r}} = \mathbf{F}$, or in our new notation

$$m \left(\frac{d^2 \mathbf{r}}{dt^2} \right)_S = \mathbf{F}$$

where \mathbf{F} is the net force acting on the particle. Combining this with our last result, multiplying through by m and rearranging produces

$$m \left(\frac{d^2 \mathbf{r}}{dt^2} \right)'_S = \mathbf{F} - 2m\boldsymbol{\omega} \times \left(\frac{d\mathbf{r}}{dt} \right)_{S'} - m\dot{\boldsymbol{\omega}} \times \mathbf{r} - m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r})$$

This is Newton's second in rotating frames of reference. To explain the particle's motion, an observer rotating with S' must introduce the existence of three further terms on the right side of the equation besides the traditional net force \mathbf{F} . The three extra terms are called *fictitious forces* and are a legitimate physical consequence of the frame S' 's rotation. Put simply, they arise because S' is not inertial. We will go over each in the next section.

Note that from the perspective of an observer rotating with S' a free particle (for which $\mathbf{F} = 0$) does not travel with uniform motion $m \left(\frac{d^2 \mathbf{r}}{dt^2} \right)'_S = 0$ because of the three fictitious forces.

1.4.3 Fictitious Forces

We will now analyze each term in the generalized second law

$$m \left(\frac{d^2 \mathbf{r}}{dt^2} \right)'_S = \mathbf{F} - 2m\boldsymbol{\omega} \times \left(\frac{d\mathbf{r}}{dt} \right)_{S'} - m\dot{\boldsymbol{\omega}} \times \mathbf{r} - m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r})$$

1. The term $-2m\boldsymbol{\omega} \times \left(\frac{d\mathbf{r}}{dt} \right)_{S'}$ is called the *Coriolis force*.

- It depends on the velocity $\left(\frac{d\mathbf{r}}{dt} \right)_{S'}$ of the particle as measured by an observer rotating with S' . Because the force is velocity dependent it is felt only by moving particles. Note also that the force is independent of position.
- The Coriolis force has an interesting effect: it makes moving particles turn in circles. In fact, you probably (sort of) already knew this: the Coriolis force is mathematically identical to the familiar Lorentz force felt by a charged particle moving in a magnetic field. Just compare the equations:

$$\mathbf{F}_{\text{cor}} \sim m\mathbf{v} \times \boldsymbol{\omega} \quad \text{and} \quad \mathbf{F}_{\text{mag}} \sim q\mathbf{v} \times \mathbf{B}$$

We know the Lorentz force makes moving charged particles turn in circles and can expect the Coriolis force to do the same. More formally, the circles result because the force is always perpendicular to velocity by virtue of the cross product $\mathbf{v} \times \boldsymbol{\omega}$.

- The magnitude of the Coriolis force is

$$|\mathbf{F}_{\text{cor}}| = 2m\omega v \sin \theta$$

where θ is the angle between the particle's velocity \mathbf{v} and the axis of rotation $\boldsymbol{\omega}$.

2. The term $-m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r})$ is called the *centrifugal force*.

- With the vector identity $\mathbf{a} \times (\mathbf{b} \times \mathbf{c}) = (\mathbf{a} \cdot \mathbf{c})\mathbf{b} - (\mathbf{a} \cdot \mathbf{b})\mathbf{c}$ it can be written

$$\begin{aligned} \mathbf{F}_{\text{cent}} &= -m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r}) \\ &= -m(\boldsymbol{\omega} \cdot \mathbf{r})\boldsymbol{\omega} + m\omega^2 \mathbf{r} \end{aligned}$$

The centrifugal force on a particle points from the particle perpendicularly away from the axis of rotation.

- The magnitude of the centrifugal force is

$$|\mathbf{F}_{\text{cent}}| = m\omega^2 \text{proj}_{\boldsymbol{\omega}}(\mathbf{r})$$

where $\text{proj}_{\boldsymbol{\omega}}(\mathbf{r})$ is the orthogonal projection of the position vector onto the axis of rotation $\boldsymbol{\omega}$.

- There's an interesting interpretation of the centrifugal force. First, note that the force does not depend on the particle's velocity, only on its position \mathbf{r} . This suggests we can view the centrifugal force as a conservative force, with an associated potential, the form

$$\mathbf{F}_{\text{cent}} = -\nabla V(\mathbf{r}) \quad \text{where} \quad V(\mathbf{r}) = -\frac{m}{2} |\boldsymbol{\omega} \times \mathbf{r}|^2$$

In the rotating frame V has the interpretation of a potential energy depending on the particle's position associated with the frame's rotation. The potential is negative. We know particles tend to minimize their energy, so we expect particles in a centrifugal potential to fly out from the axis of rotation; increasing $|r|$ will minimize the potential energy. This agrees with our earlier conclusion that that centrifugal force acts outward from the axis of rotation.

3. The term $-m\dot{\boldsymbol{\omega}} \times \mathbf{r}$ is called the *Euler force*. It arises only when S' rotates with variable angular velocity $\boldsymbol{\omega}$ (i.e. $\dot{\boldsymbol{\omega}} \neq 0$). Because most everyday examples of non-inertial frames rotate with (nearly) constant angular velocity, the Euler force is usually negligible.

1.4.4 Rotation With Moving Origins

Staying with our frames S and S' , we now consider the general case when the origins of S and S' are separated by the vector \mathbf{R} , and the origins move relative to each other in space.

S still rotates with angular velocity $\boldsymbol{\omega}$ about an axis through its center.

A real-world example would be an (almost) inertial system S at the center of the sun and a non-inertial system S' on the Earth; the S' Earth system both orbits around the sun (relative motion of origins) and rotates about an axis through its center.

We have

$$\begin{aligned}\mathbf{r} &= \mathbf{r}_O + \mathbf{r}' \\ \mathbf{v}_L &= \mathbf{v}_M + \mathbf{v}_O + \boldsymbol{\omega} \times \mathbf{r}' \\ \mathbf{a}_L &= \mathbf{a}_M + \mathbf{a}_O + \dot{\boldsymbol{\omega}} \times \mathbf{r}' + 2\boldsymbol{\omega} \times \mathbf{v}_M + \boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r}') \\ m\mathbf{a}_M &= \mathbf{F} - m\mathbf{a}_O - m\dot{\boldsymbol{\omega}} \times \mathbf{r}' - 2m\boldsymbol{\omega} \times \mathbf{v}_M - m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r}')\end{aligned}$$

where \mathbf{v}_O and \mathbf{a}_O are the relative velocity and accelerations of the origins. They are measured in the inertial frame L .

Or, using different symbols:

The dynamics of rotating coordinate system M whose origin is separated from the origin of a lab system L by the vector \mathbf{R} , including fictitious forces, are governed by the equations

$$\begin{aligned}\mathbf{r} &= \mathbf{R} + \mathbf{r}' \\ \mathbf{v}_L &= \mathbf{v}_M + \dot{\mathbf{R}}_L + \boldsymbol{\omega} \times \mathbf{r}' \\ \mathbf{a}_L &= \mathbf{a}_M + \ddot{\mathbf{R}}_L + \dot{\boldsymbol{\omega}} \times \mathbf{r}' + 2\boldsymbol{\omega} \times \mathbf{v}_M + \boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r}') \\ m\mathbf{a}_M &= \mathbf{F} - m\ddot{\mathbf{R}}_L - m\dot{\boldsymbol{\omega}} \times \mathbf{r}' - 2m\boldsymbol{\omega} \times \mathbf{v}_M - m\boldsymbol{\omega} \times (\boldsymbol{\omega} \times \mathbf{r}')\end{aligned}$$

2 Lagrangian Mechanics

2.1 Lagrange's Equations and the Principle of the Least Action

2.1.1 Coordinates and Degrees of Freedom

Lagrangian mechanics uses slightly different coordinates than the Newtonian position vector \mathbf{r} , so let's start by sorting our coordinates out. Suppose we are dealing with a system of N particles under the influence of a conservative net force.

- We need $3N$ coordinates to uniquely describe a system of N particles in 3 dimensions. N is for the number of particles in the system and 3 because each particle can move in three dimensions.

In Newtonian mechanics we usually package these coordinates into position vectors $\mathbf{r}_1, \mathbf{r}_2, \dots, \mathbf{r}_N$ where $\mathbf{r}_i = (x_i, y_i, z_i)$ are the coordinates of the i th mass point and $i = 1, \dots, N$.

- Some terminology: a three dimensional system of N particles has $3N$ *degrees of freedom*; each of the N particles contributes three. Degrees of freedom tell us how many quantities we must know to uniquely specify the configuration (position) of our system in space. For example, we need three coordinates to uniquely describe the position of a point particle in 3 dimensions, so a single particle has 3 degrees of freedom.

Our system's $3N$ degrees of freedom parameterize a $3N$ dimensional space called the *configuration space*. Each $3N$ -dimensional point in the phase space represents a configuration of the system (i.e. the position of all N particles). The time evolution of our system corresponds to a $3N$ -dimensional curve in configuration space.

- Our new coordinates: We will replace the N Newton-style position vectors \mathbf{r}_i with $3N$ coordinates x^A where $A = 1, \dots, 3N$. Note that the A is not a power but an "up" index; you'll see more of this notation in tensor calculus. For our purposes in this course the placement of the indexes is not critical; you can read x^A as x_A .
- On to Newton's laws with our new coordinates x^A . For our conservative system we have $\mathbf{F} = -\nabla V$. In Newtonian mechanics the equation of motion for the i th particle would be $\dot{\mathbf{p}}_i = -\frac{\partial V}{\partial \mathbf{r}_i}$. With our new coordinates, Newton's equations of motion read

$$\dot{p}_A = -\frac{\partial V}{\partial x^A}, \quad A = 1, \dots, 3N$$

where $p_A = m_A \dot{x}^A$.

2.1.2 Lagrangian and Principle of Least Action

We continue with our system of $3N$ particles under the influence of a conservative force. We start by defining the Lagrangian function, the cornerstone of Lagrangian mechanics.

- Our system's *Lagrangian* function, or simply *Lagrangian*, is a function of the positions x^A and velocities \dot{x}^A given by

$$L(x^A, \dot{x}^A) = T(\dot{x}^A) - V(x^A)$$

where $T = \frac{1}{2} \sum_A m_A (\dot{x}^A)^2$ is the system's total kinetic energy and $V(x^A)$ is the system's total potential energy.

- Recall the time evolution of our system corresponds to a $3N$ dimensional path in configuration space. An *action* is a property of a path $x(t)$ in configuration space, and is defined as

$$S[x(t)] = \int_{t_i}^{t_f} L(x(t), \dot{x}(t)) dt$$

where $L(x, \dot{x}) = T(\dot{x}) - V(x)$ is the Lagrangian.

- For the principle of least action, we consider all smooth paths in the configuration space between the fixed endpoints x_i and x_f where $x(t_i) = x_i$ and $x(t_f) = x_f$.

The *principle of least action states*:

The path taken by a system between two points in its configuration space is an extremum of the action S .

The extremum could be a maximum, minimum, or saddle point.

In fact, all the fundamental laws of physics can be written in terms of an action principle, but that's another story.

2.1.3 Proof of Least Action Principle

- Consider a smooth paths through configuration space between the fixed endpoints x_i^A and x_f^A where $x^A(t_i) = x_i^A$ and $x^A(t_f) = x_f^A$. We'll write the path $x^A(t)$. Then consider varying the path just slightly. Instead of $x^A(t)$ we have

$$x^A \rightarrow x^A(t) + \delta x^A(t)$$

where δx^A is small compared to x^A .

- Although the interior of the path can vary as $x^A \rightarrow x^A(t) + \delta x^A(t)$, we'll insist on keeping our endpoints fixed, so $x^A(t_i) = x_i^A$ and $x^A(t_f) = x_f^A$. In other words, the variation vanishes at the endpoints x_i^A and x_f^A and

$$\delta x^A(t_i) = \delta x^A(t_f) = 0 \implies x^A(t_i) = x_i^A \quad \text{and} \quad x^A(t_f) = x_f^A$$

If the talk of variation of paths in configuration space seems abstract, think in everyday terms of walking between the same two points x_i and x_f , but taking a slightly different path between them.

- The change in action δS corresponding to our slight change in path $\delta x^A(t)$ is

$$\begin{aligned} \delta S &= \delta \left[\int_{t_i}^{t_f} L(x^A, \dot{x}^A) dt \right] = \int_{t_i}^{t_f} \delta L(x^A, \dot{x}^A) dt \\ &= \int_{t_i}^{t_f} \left(\frac{\partial L}{\partial x^A} \delta x^A + \frac{\partial L}{\partial \dot{x}^A} \delta \dot{x}^A \right) dt \\ &= \int_{t_i}^{t_f} \frac{\partial L}{\partial x^A} \delta x^A dt + \int_{t_i}^{t_f} \frac{\partial L}{\partial \dot{x}^A} \delta \dot{x}^A dt \end{aligned}$$

- We integrate the second integral by parts with

$$dv = \delta \dot{x}^A dt \implies v = \delta x^A \quad \text{and} \quad u = \frac{\partial L}{\partial \dot{x}^A} \implies du = \frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A}$$

Applying $\int u dv = uv - \int v du$ and joining the two integrals into one integrand, the variation in action reads

$$\delta S = \int_{t_i}^{t_f} \left[\frac{\partial L}{\partial x^A} - \frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A} \right] \delta x^A dt + \left[\frac{\partial L}{\partial \dot{x}^A} \delta x^A \right]_{t_i}^{t_f}$$

Since we have fixed the endpoints $\delta x^A(t_i) = \delta x^A(t_f) = 0$ the last term (the uv term from integration by parts) vanishes. We're left with

$$\delta S = \int_{t_i}^{t_f} \left[\frac{\partial L}{\partial x^A} - \frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A} \right] \delta x^A dt$$

- Back to the principle of least action: the requirement that action S is an extremum holds only if $\delta S = 0$ for all small variations in the path $\delta x^A(t)$. Since $\delta x^A \neq 0$, the requirement $\delta S = 0$ for all paths holds if the other term in the integrand is zero, namely

$$\frac{\partial L}{\partial x^A} - \frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A} = 0 \quad \text{for all } A = 1, \dots, 3N$$

These equations are called *Lagrange's equations*, and we'll see they replace Newton's equations in Lagrangian mechanics.

- To finish the proof of the least action principle, we just need to show that Lagrange's equations are equivalent to Newton's. From the definition of the Lagrangian $L(x^A, \dot{x}^A) = T(\dot{x}^A) - V(x^A)$, we have

$$\frac{\partial L}{\partial x^A} = -\frac{\partial V}{\partial x^A} \quad \text{and} \quad \frac{\partial L}{\partial \dot{x}^A} = \frac{\partial T}{\partial \dot{x}^A} = \frac{\partial}{\partial \dot{x}^A} \left[\frac{1}{2} m (\dot{x}^A)^2 \right] = p_A$$

Putting the pieces together, we see the Lagrange and Newton equations are equivalent

$$\frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A} = \frac{\partial L}{\partial x^A} \iff \dot{p}_A = -\frac{\partial V}{\partial x^A} \iff \dot{\mathbf{p}}_i = \mathbf{F}_i$$

Since Newton's equations are known to hold, so must Lagrange's, meaning $\frac{\partial L}{\partial x^A} - \frac{d}{dt} \frac{\partial L}{\partial \dot{x}^A} = 0$, which means $\delta S = 0$ for all variations $\delta x^A(t)$. This concludes the proof of the least action principle.

2.1.4 Lagrange Equations In Changing Coordinate Systems

Lagrangian mechanics have two strong advantages over Newtonian mechanics: Lagrange's equations hold in any coordinate system (Newton's second law holds only in inertial systems) and the Lagrangian formalism greatly simplifies systems with constraints. We'll start by examining Lagrange's equations in changing coordinate systems and deal with constraints in the next section.

- **Notation:** First, a note on a wonderfully efficient notation we'll use for the rest of the course, called the *Einstein summation convention*. When an index appears twice in a single term without further specification, it implies summation of that term over all values of the index. Here's a simple example using the index $i = 1, 2, 3$:

$$a_i b^i \equiv \sum_{i=1}^3 a_i b_i = a_1 b_1 + a_2 b_2 + a_3 b_3$$

Recognize this? It's the dot product of the vectors $\mathbf{a} \cdot \mathbf{b}$.

Here's a concrete example of how we'll use the convention in this course. Consider some quantity q (they'll turn out to be coordinates) that is a function of the coordinates $x_A, A = 1, \dots, 3N$ and time, so $q = q(x_1, \dots, x_{3N}, t)$. Using the chain rule, summation convention and dot notation for time derivatives, the time derivative of q is then

$$\dot{q} = \frac{\partial q}{\partial x^A} \dot{x}^A + \frac{\partial q}{\partial t} \quad \text{or, equivalently} \quad \frac{dq}{dt} = \sum_{A=1}^{3N} \frac{\partial q}{\partial x_A} \frac{dx^A}{dt} + \frac{\partial q}{\partial t}$$

Once you get a feel for it, the left hand side is a lot cleaner.

- With summation convention behind us, we turn our focus to Lagrange's equations in changing coordinate systems. Staying with our system of N particles with coordinates x_A , let's introduce a new coordinate system

$$q_a = q_a(x_1, \dots, x_{3N}, t)$$

where $a = 1, \dots, 3N$ and the t dependence allows for a coordinate system that changes with time. As already written earlier, q_a change with time as

$$\dot{q} = \frac{\partial q}{\partial x^A} \dot{x}^A + \frac{\partial q}{\partial t}$$

- For the q_a to be a well-defined coordinate system we should be able to invert the relationship to recover the original x^A as $x^A = x^A(q_1, \dots, q_{3N}, t)$. Formally, this is possible as long as the $3N \times 3N$ matrix $\frac{\partial x^A}{\partial q_a}$ is non-singular, i.e.

$$\det \frac{\partial x^A}{\partial q_a} \neq 0$$

but this condition is more a mathematical technicality than something we have to deal with in real systems.

- Assuming we can invert $q_a(x^A, t)$ to recover $x^A = x^A(q_a, t)$, we have the time derivative

$$\dot{x}^A = \frac{\partial x^A}{\partial q_a} \dot{q}_a + \frac{\partial x^A}{\partial t}$$

Armed with expressions for x^A and \dot{x}^A in terms of the new coordinates q_a , we can write the Lagrangian $L(x^A, \dot{x}^A)$ in terms of q_a :

$$L(x^A, \dot{x}^A) = L(x^A(q_a, t), \dot{x}^A(q_a, \dot{q}_a, t))$$

We now have a bit of derivative work to do: we need to calculate $\frac{\partial L}{\partial \dot{q}_a}$ and $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_a} \right)$. The chain rule tells us

$$\frac{\partial L}{\partial q_a} = \frac{\partial L}{\partial x^A} \frac{\partial x^A}{\partial q_a} + \frac{\partial L}{\partial \dot{x}^A} \left(\frac{\partial^2 x^A}{\partial q_a \partial q_b} \dot{q}_b + \frac{\partial^2 x^A}{\partial q_a \partial t} \right)$$

and

$$\frac{\partial L}{\partial \dot{q}_a} = \frac{\partial L}{\partial \dot{x}^A} \frac{\partial \dot{x}^A}{\partial \dot{q}_a}$$

To simplify $\frac{\partial L}{\partial \dot{q}_a}$ we'll need the intermediate result $\frac{\partial \dot{x}^A}{\partial \dot{q}_a} = \frac{\partial x^A}{\partial q_a}$, which we find by taking the partial derivative of \dot{x}^A with respect to \dot{q}_a :

$$\frac{\partial \dot{x}^A}{\partial \dot{q}_a} = \frac{\partial}{\partial \dot{q}_a} \left[\frac{\partial x^A}{\partial q_a} \dot{q}_a \right] + \frac{\partial}{\partial \dot{q}_a} \left[\frac{\partial x^A}{\partial t} \right] = 0 + \frac{\partial x^A}{\partial q_a} \frac{\partial \dot{q}_a}{\partial \dot{q}_a} + 0 \implies \frac{\partial \dot{x}^A}{\partial \dot{q}_a} = \frac{\partial x^A}{\partial q_a}$$

With this result we have

$$\frac{\partial L}{\partial \dot{q}_a} = \frac{\partial L}{\partial \dot{x}^A} \frac{\partial \dot{x}^A}{\partial \dot{q}_a} = \frac{\partial L}{\partial \dot{x}^A} \frac{\partial x^A}{\partial q_a}$$

The time derivative of $\frac{\partial L}{\partial \dot{q}_a} = \frac{\partial L}{\partial \dot{x}^A} \frac{\partial x^A}{\partial q_a}$, using the product and chain rules, is

$$\begin{aligned} \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_a} \right) &= \frac{d}{dt} \left[\frac{\partial L}{\partial \dot{x}^A} \right] \frac{\partial x^A}{\partial q_a} + \frac{d}{dt} \left[\frac{\partial x^A}{\partial q_a} \right] \frac{\partial L}{\partial \dot{x}^A} \\ &= \frac{d}{dt} \left[\frac{\partial L}{\partial \dot{x}^A} \right] \frac{\partial x^A}{\partial q_a} + \frac{\partial L}{\partial \dot{x}^A} \left(\frac{\partial^2 x^A}{\partial q_a \partial q_b} \dot{q}_b + \frac{\partial^2 x^A}{\partial q_a \partial t} \right) \end{aligned}$$

- We now combine the equations for $\frac{\partial L}{\partial \dot{q}_a}$ and $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_a} \right)$. The term $\frac{\partial L}{\partial \dot{x}^A} \left(\frac{\partial^2 x^A}{\partial q_a \partial q_b} \dot{q}_b + \frac{\partial^2 x^A}{\partial q_a \partial t} \right)$ conveniently cancels. Grouping the remaining $\frac{\partial x^A}{\partial q_a}$ terms gives

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_a} \right) - \frac{\partial L}{\partial \dot{q}_a} = \left[\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{x}^A} \right) - \frac{\partial L}{\partial x^A} \right] \frac{\partial x^A}{\partial q_a}$$

This is an excellent result: it tells us that if Lagrange's equations hold in the x^A coordinate system they also hold in the q_a system. In other words, Lagrange's equations are preserved across any coordinate transformation of the form $q_a = q_a(x_1, \dots, x_{3N}, t)$.

How to see this result? If Lagrange's equations hold in the x^A coordinates then by definition of the Lagrange equation $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{x}^A} \right) - \frac{\partial L}{\partial x^A} = 0$. But if $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{x}^A} \right) - \frac{\partial L}{\partial x^A} = 0$ then our result tells us that $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_a} \right) - \frac{\partial L}{\partial \dot{q}_a} = 0$, which means Lagrange's equations also hold in the q_a coordinates. The converse (i.e. the LE holding for q_a implying the LE hold for x_A) holds as long as $\det \frac{\partial x^A}{\partial q_a} \neq 0$.

2.2 Constraints and Generalized Coordinates

We continue with our system of N particles and coordinates $x_A, A = 1, \dots, 3N$. So far we have tacitly assumed the particles in our system can roam anywhere in \mathbb{R}^3 . But what if there are constraints restricted where the particles can move? In Newton mechanics we handle constraints with "constraint forces", like the normal force of a surface preventing an object from falling through the floor. Lagrangian mechanics greatly simplify the treatment of constraints. But first, some background.

2.2.1 Holonomic Constraints

Formally, constraints are conditions that reduce a system's degrees of freedom. In Lagrangian mechanics constraints are often desirable since they make a system's motion simpler.

- A *holonomic constraint* is a relationship between the coordinates x^A that can be written in the form

$$f_\alpha(x_A, t) = 0 \quad \alpha = 1, \dots, 3N - n$$

where t allows for the constraint to depend on time. A system as a whole is called a *holonomic system* if all of its constraints are holonomic.

- Holonomic constraints can be solved in terms of n *generalized coordinates* q_1, \dots, q_n to which we will assign the index i and label q_i . Once we solve a system's constraints in terms of the generalized coordinates q_i we can transform from the original coordinates x^A to the generalized coordinates q_i :

$$x_A = x_A(q_1, \dots, q_n)$$

A system with $3N - n$ constraints and thus n generalized coordinates has n degrees of freedom, since we need n generalized coordinates to uniquely describe the system.

- Since all the indexes and variables can be confusing, here's a quick review. Feel free to skip. To have a simple system as possible, we want as many constraints as possible. A completely free system has zero constraints and a completely fixed system that can't move at all has $3N$ constraints. In our notation the system has $3N - n$ constraints, so n should be small to maximize the number of constraints. If n is small, we only need a small number of generalized coordinates to uniquely describe our system. This is good, we want to be able to describe the system with as few coordinates as possible.
- Here's how Lagrangian mechanics handles constraints. For each of the $3N - n$ constraints f_α we introduce $3N - n$ new variables λ_α , called *Lagrange multipliers*.

In terms of the Lagrange multipliers we define a new Lagrangian

$$L' = L(x^A, \dot{x}^A) - \lambda_\alpha f_\alpha(x^A, t)$$

We can treat the Lagrange multipliers λ_α as new coordinates. Using our new Lagrangian L' the Lagrange-Euler equations for λ_α are then

$$\frac{\partial L'}{\partial \lambda_\alpha} - \frac{d}{dt} \left(\frac{\partial L'}{\partial \dot{\lambda}_\alpha} \right) = 0 \implies f_\alpha(x^A, t) = 0$$

This simply recovers the constraints $f_\alpha(x^A, t) = 0$.

The Lagrangian equations for x^A are more interesting. We have

$$\frac{\partial L'}{\partial x^A} - \frac{d}{dt} \left(\frac{\partial L'}{\partial \dot{x}^A} \right) = 0 \implies \frac{\partial L}{\partial x^A} + \lambda_\alpha \frac{\partial f_\alpha}{\partial x^A} - \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{x}^A} \right) = 0$$

This result will tell us how the Lagrangian formalism handles constraints.

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{x}^A} \right) - \frac{\partial L}{\partial x^A} = \lambda_\alpha \frac{\partial f_\alpha}{\partial x^A}$$

Notice the left hand side is the familiar Lagrange equation of motion for the unconstrained system. The right hand side is the manifestation of the constraint forces in our system. In principle, these equations could be solved just like in the Newtonian formalism. But there's a much easier way to go about it using the generalized coordinates q_i .

2.2.2 Holonomic Constraints and Generalized Coordinates

- The Lagrangian formalism makes it possible to solve for the motion of a constrained system in terms of the n generalized coordinates q_i instead of $3N$ coordinates x^A .

This is really useful. Why? Because n is usually small and $3N$ is usually large.¹ It's easier to solve a few equations instead of many equations.

- Derive the system's equations of motion directly in the generalized coordinates means we must be able to write the Lagrangian in terms of the q_i in the form:

$$L(q_i, \dot{q}_i, t) = L(x^A(q_i, t), \dot{x}^A(q_i, \dot{q}_i, t))$$

following is a quick proof.

- We'll work with the modified Lagrangian

$$L' = L(x^A, \dot{x}^A) - \lambda_\alpha f_\alpha(x^A, t)$$

and change from the x^A coordinates to the generalized coordinates and Lagrange multipliers:

$$x^A \rightarrow \begin{cases} q_i, & i = 1, \dots, n \\ \lambda_\alpha, & \alpha = 1, \dots, 3N \end{cases}$$

Notice we start with $3N$ coordinates x^A and end with $n + 3N - n = 3N$ coordinates, so our accounting is okay.

- Using our new Lagrangian $L' = L'(q_i, \dot{q}_i, \lambda_\alpha, t)$ we can find the Lagrange equations for q_i (just like in the previous section 2.2.1 with $L'(x^A, \dot{x}^A, \lambda_\alpha, t)$). They are

$$\frac{\partial L'}{\partial q_i} - \frac{d}{dt} \left(\frac{\partial L'}{\partial \dot{q}_i} \right) = 0 \implies \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} = \lambda_\alpha \frac{\partial f_\alpha}{\partial q_i}$$

But because the constraints f_α are by definition independent of the generalized coordinates q_i , we have $\lambda_\alpha \frac{\partial f_\alpha}{\partial q_i} = 0$ leaving

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} = 0$$

This is a great result. It tells us that if we're interested only in the dynamics of the generalized coordinates q_i we can work directly with the Lagrangian $L(q_i, \dot{q}_i, t)$ and ignore the Lagrange multipliers. This basically means we only have to solve n equations instead of $3N$. If our system is highly constrained, $3N - n$ is a big difference!

2.2.3 A Summary of Constraints and Coordinates in Lagrangian Mechanics

Keeping with a system of N particles in three dimensions.

- In Lagrangian mechanics we begin with $3N$ coordinates x^A ; $3N$ is the default number of coordinates needed to describe N particles in 3 dimensions.

Usually the system has some constraints. These simplify the motion and allow us to describe the dynamics with fewer coordinates, called generalized coordinates.

¹At least in introductory classical dynamics courses :)

For each additional constraint, we need one fewer coordinate to describe our system. If we have $3N - n$ constraints, we then need only $3N - (3N - n) = n$ generalized coordinates. Basically, the generalized coordinates are the simplified coordinates left over after applying all the constraints.

- Because it's fewer variable to work with, we describe a system in terms of the n generalized coordinates q_i with $i = 1, \dots, n$. These coordinates form an n -dimensional space called the configuration space.

The system's time evolution corresponds to a n -dimensional curve in configuration space. The time evolution is governed by the Lagrangian

$$L = L(q_i, \dot{q}_i, t) = T(\dot{q}_i) - V(q_i)$$

such that the coordinates q_i obey the Lagrange equations

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} = 0$$

where a system of n second order (generally) non-linear differential equations. The goal of Lagrangian mechanics is then solving the Lagrange equations for the dynamics of the generalized coordinates q_i .

- Finally, for the sake of completeness, a few words on something we won't have to deal with in this course. But it's good to know such things exist.

Constraints that cannot be written in the form $f(x_A, t) = 0$ are called *non-holonomic* constraints. Two common examples are: constraints with inequalities and velocity-dependent constraints that cannot be integrated into the form $f(x^A, t) = 0$, e.g;

$$\begin{aligned} \text{Constraint with an inequality:} & \quad f(x_A, t) \leq C, \quad (C \in \mathbb{R}) \\ \text{Velocity dependent constraint:} & \quad g(x^A, \dot{x}^A, t) = 0 \end{aligned}$$

And a some vocabulary: If a constraint has an explicit time dependence, it is called *rheonomic*. If a constraint is not explicitly time-dependent, it is called *scleronomic*.

2.3 Lagrange's Equations and D'Alembert's Principle

D'Alembert's principle is another way to derive Lagrange's equations using the concept of virtual displacements and virtual work. By the way, even though we started with the principle of least action, derivation by D'Alembert's principle is the conventional way to introduce Lagrange's equations. I think this version is more work, so I put the principle of least action derivation first.

2.3.1 Virtual Displacements and D'Alembert's Principle

- First, we split the forces acting on our system into active forces and constraint forces.

$$\mathbf{F}_i = \mathbf{F}_i^a + \mathbf{F}_i^c$$

Active forces are the forces that directly affect the dynamics of the generalized coordinates. The constraints, well, constrain the possible paths our system can take. They're the forces associated with the Lagrange multipliers, which we've shown we can ignore if we're interested only in the dynamics of the generalized coordinates. In

general, constraint forces are perpendicular to a system's motion (think normal force, tension forces, etc...), so they can't do work on the system.

We'll come back to the separation of forces $\mathbf{F}_i = \mathbf{F}_i^a + \mathbf{F}_i^c$ shortly.

- Constraint forces that are perpendicular to the system's motion for all displacements, including virtual displacements, are called *ideal constraints*. These are the common constraints like normal force and tension. For this section and this course, we'll consider only ideal constraints. For a system with ideal constraints, in which the constraint forces are always perpendicular to motion, the work done by constraint forces is zero.
- We'll return to our vector notation, using N position vectors $\mathbf{r}_i, i = 1, \dots, N$ instead of the $3N$ Lagrange-style coordinates x^A .

A definition: a *virtual displacement* of a system is a change in the configuration of the system as the result of any arbitrary infinitesimal change of the coordinates *at a given instant in time* consistent with the forces and constraints imposed on the system at that instant in time.

The displacement is called virtual to distinguish it from an actual displacement in a time *interval* dt , during which the the forces and constraints might change.

Think of a virtual displacement as a sort of freeze-frame, capturing the system's possible displacement at a single moment in time for given the forces and constraints acting at that instant.

To distinguish them from displacements $d\mathbf{r}$ than occur in a finite time interval dt , we denote virtual displacements by $\delta\mathbf{r}$.

Virtual work is the work of a force in the direction of a virtual displacement, $\delta W = \mathbf{F} \cdot \delta\mathbf{r}$.

- First, we'll consider a static equilibrium situation, then move to dynamics. For a system in equilibrium, the net force on each particle vanishes, so $\mathbf{F}_i = 0$. If $\mathbf{F}_i = 0$ then naturally any work done by the net force also vanishes. We then have

$$\mathbf{F}_i^{(\text{eq})} \cdot \delta\mathbf{r}_i = 0$$

where the notation stresses that the net force $\mathbf{F}_i^{(\text{eq})}$ is such that the system is in equilibrium. Using the general decomposition of the net force into $\mathbf{F}_i^{(\text{eq})} = \mathbf{F}_i^{(a)} + \mathbf{F}_i^{(c)}$, this becomes

$$\left(\mathbf{F}_i^{(a)} + \mathbf{F}_i^{(c)} \right) \cdot \delta\mathbf{r}_i = 0$$

Restricting ourselves to a system with ideal constraint, $\mathbf{F}_i^{(c)}$ will always be perpendicular to $\delta\mathbf{r}_i$, so the virtual work done by the constraint forces vanishes. We then have

$$\mathbf{F}_i^{(a)} \cdot \delta\mathbf{r}_i = 0$$

for any system in equilibrium. This result is called the *principle of virtual work*. It gives an condition for equilibrium for any system with ideal constraints: the total virtual work done by the active forces must vanish.

Note, however, that this condition does *not* require that the virtual work on each particle is zero, just that the sum of the individual work (i.e. the total work) adds

up to zero. This means that the *individual* $\mathbf{F}_i^{(a)}$ need not equal zero for a system in equilibrium.

- On to dynamics—we'll use little trick here. We take Newton's equation of motion $\mathbf{F}_i = \dot{\mathbf{p}}_i$, which applies to an accelerating particle, then rearrange it to get

$$\mathbf{F}_i^{(\text{eq})} \equiv \mathbf{F}_i - \dot{\mathbf{p}}_i = 0$$

which we interpret, à la Goldstein, as follows: the accelerating system would be in equilibrium if acted on by a net force $\mathbf{F}_i^{(\text{eq})}$ equal to the accelerating force \mathbf{F}_i minus a “reversed effective force” $\dot{\mathbf{p}}_i$. We can then plug $\mathbf{F}_i^{(\text{eq})} = (\mathbf{F}_i - \dot{\mathbf{p}}_i)$ into the equilibrium equation $\mathbf{F}_i^{(\text{eq})} \cdot \delta \mathbf{r}_i = 0$ to get

$$(\mathbf{F}_i - \dot{\mathbf{p}}_i) \cdot \delta \mathbf{r}_i = 0$$

Decomposing the net force into the active and constraint force $\mathbf{F}_i = \mathbf{F}_i^{(a)} + \mathbf{F}_i^{(c)}$ and then recalling that the virtual work $\mathbf{F}_i^{(c)} \cdot \delta \mathbf{r}_i$ is zero for ideal constraints, we get

$$\begin{aligned} 0 &= \left[(\mathbf{F}_i^{(a)} + \mathbf{F}_i^{(c)}) - \dot{\mathbf{p}}_i \right] \cdot \delta \mathbf{r}_i = (\mathbf{F}_i^{(a)} - \dot{\mathbf{p}}_i) \cdot \delta \mathbf{r}_i + \mathbf{F}_i^{(c)} \cdot \delta \mathbf{r}_i \\ &= (\mathbf{F}_i^{(a)} - \dot{\mathbf{p}}_i) \cdot \delta \mathbf{r}_i + 0 \end{aligned}$$

The result

$$(\mathbf{F}_i^{(a)} - \dot{\mathbf{p}}_i) \cdot \delta \mathbf{r}_i = 0$$

is called *D'Alembert's principle*. The result is important in that it applies to dynamics as well as statics, and only contains the active forces. I stress again that the quantity $\mathbf{F}_i^{(a)} - \dot{\mathbf{p}}_i$ is *not* generally zero for each particle individually, just that all the $\mathbf{F}_i^{(a)} - \dot{\mathbf{p}}_i$ for the particles together sum to zero.

2.3.2 Deriving the Lagrange Equations: Take Two, Part 1

- D'Alembert's principle (which I'm now writing with the index j —explained shortly)

$$(\mathbf{F}_j^{(a)} - \dot{\mathbf{p}}_j) \cdot \delta \mathbf{r}_j = 0$$

is an important result, but it can't quite get us to Lagrange's equations. The problem is that in the presence of constraints, the coordinates \mathbf{r}_j are in general dependent on each other.

For our equations of motion to be well defined, we'll need to introduce a set of *independent* coordinates. In fact, we have just what we need for the job: the generalized coordinates q_i . The goal will be to transform D'Alembert's principle from \mathbf{r}_j to q_i .

That's also why I changed indices: we had previously used i for the generalized coordinates q_i , and I wanted to stay consistent. We can't use i for both the vector coordinates \mathbf{r} and generalized coordinates q since there are different number of each, so we'll use j for the \mathbf{r}_j . It's just notational struggles—nothing of physical significance.

- We start with the transformation, introduced in Subsection 2.2.1,

$$\mathbf{r}_j = \mathbf{r}_j(q_1, \dots, q_n, t)$$

The total derivative of \mathbf{r}_j is then

$$d\mathbf{r}_j = \frac{\partial \mathbf{r}_j}{\partial q_i} dq_i + \frac{\partial \mathbf{r}_j}{\partial t} dt$$

For virtual displacements, which by definition occur for $dt = 0$, we have

$$\delta \mathbf{r}_j = \frac{\partial \mathbf{r}_j}{\partial q_i} \delta q_i$$

While the total virtual work $\mathbf{F}_j \cdot \delta \mathbf{r}_j$ becomes

$$\mathbf{F}_j \cdot \delta \mathbf{r}_j = \mathbf{F}_j \cdot \left(\frac{\partial \mathbf{r}_j}{\partial q_i} \delta q_i \right) \equiv Q_i \delta q_i$$

where we defined a *generalized force* with components

$$Q_i = \mathbf{F}_j \cdot \frac{\partial \mathbf{r}_j}{\partial q_i}$$

The transformation $\mathbf{F}_j \cdot \delta \mathbf{r}_j = Q_i \delta q_i$ takes care of the first term $\mathbf{F}_j^{(a)} \cdot \delta \mathbf{r}_j$ in D'Alembert's principle, where we could safely equate $\mathbf{F}_j \cdot \delta \mathbf{r}_j$ to $\mathbf{F}_j^{(a)} \cdot \delta \mathbf{r}_j$ because D'Alembert's principle requires a priori that the constraint forces do no work, so the $\mathbf{F}^{(c)}$ terms don't contribute to the virtual work $Q_i \delta q_i$.

- The transformation from \mathbf{r}_j to q_i in the second term $\dot{\mathbf{p}}_j \cdot \delta \mathbf{r}_j$ requires a bit more effort. Writing $\dot{\mathbf{p}} = m\ddot{\mathbf{r}}$ and substituting the expression for $\delta \mathbf{r}_j(q_i)$, we get

$$\dot{\mathbf{p}}_j \cdot \delta \mathbf{r}_j = m_j \ddot{\mathbf{r}}_j \cdot \left(\frac{\partial \mathbf{r}_j}{\partial q_i} \delta q_i \right)$$

Next, we'll use a little trick. By reverse engineering the product rule, we can write

$$m_j \ddot{\mathbf{r}}_j \frac{\partial \mathbf{r}_j}{\partial q_i} = \frac{d}{dt} \left[m_j \dot{\mathbf{r}}_j \frac{\partial \mathbf{r}_j}{\partial q_i} \right] - \dot{\mathbf{r}}_j \frac{d}{dt} \left[m_j \frac{\partial \mathbf{r}_j}{\partial q_i} \right]$$

Getting an expression for $\dot{\mathbf{r}}_j$ isn't too hard: a time differentiation of $\mathbf{r}_j = \mathbf{r}_j(q_1, \dots, q_n)$ with the chain rule gives

$$\dot{\mathbf{r}}_j = \frac{d\mathbf{r}_j}{dt} = \frac{\partial \mathbf{r}_j}{\partial q_i} \dot{q}_i + \frac{\partial \mathbf{r}_j}{\partial t}$$

- The next step is basically the foundation of the Lagrangian formalism (i.e. it's important): we define the \dot{q}_i as a *new variable, independent of the q_i* . Even though they both involve the coordinates q , the Lagrangian formalism treats the q_i and \dot{q}_i as if they were just as different as the hypothetical variables x and y .

If we can view the \dot{q}_i as new, independent coordinates, we can take the partial derivative of $\dot{\mathbf{r}}_j = \frac{\partial \mathbf{r}_j}{\partial q_i} \dot{q}_i + \frac{\partial \mathbf{r}_j}{\partial t}$ with respect to \dot{q}_i with the product rule to get

$$\begin{aligned} \frac{\partial \dot{\mathbf{r}}_j}{\partial \dot{q}_i} &= \frac{\partial}{\partial \dot{q}_i} \left[\frac{\partial \mathbf{r}_j}{\partial q_i} \right] \dot{q}_i + \frac{\partial \mathbf{r}_j}{\partial q_i} \frac{\partial \dot{q}_i}{\partial \dot{q}_i} + \frac{\partial}{\partial \dot{q}_i} \left[\frac{\partial \mathbf{r}_j}{\partial t} \right] \\ &= 0 + \frac{\partial \mathbf{r}_j}{\partial q_i} + 0 \implies \frac{\partial \dot{\mathbf{r}}_j}{\partial \dot{q}_i} = \frac{\partial \mathbf{r}_j}{\partial q_i} \end{aligned}$$

The result $\frac{\partial \dot{\mathbf{r}}_j}{\partial \dot{q}_i} = \frac{\partial \mathbf{r}_j}{\partial q_i}$ looks like we just trivially "canceled the dots" but we could only do that because we defined the \dot{q}_i as independent coordinates.

- Back to where we left off with

$$m_j \ddot{\mathbf{r}}_j \frac{\partial \mathbf{r}_j}{\partial q_i} = \frac{d}{dt} \left[m_j \dot{\mathbf{r}}_j \frac{\partial \mathbf{r}_j}{\partial q_i} \right] - \dot{\mathbf{r}}_j \frac{d}{dt} \left[m_j \frac{\partial \mathbf{r}_j}{\partial q_i} \right] \equiv \text{I} - \text{II}$$

where we'll tackle the expression term by term. Starting with I, we'll use the result $\frac{\partial \dot{\mathbf{r}}_j}{\partial \dot{q}_i} = \frac{\partial \mathbf{r}_j}{\partial q_i}$ and a reverse engineered chain rule to get

$$\text{I} \equiv \frac{d}{dt} \left[m_j \dot{\mathbf{r}}_j \frac{\partial \mathbf{r}_j}{\partial q_i} \right] = \frac{d}{dt} \left[m_j \dot{\mathbf{r}}_j \frac{\partial \dot{\mathbf{r}}_j}{\partial \dot{q}_i} \right] = \frac{d}{dt} \left[\frac{\partial}{\partial \dot{q}_i} \left(\frac{m_j \dot{\mathbf{r}}_j^2}{2} \right) \right] = \frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right)$$

Look who showed up: it's the system's total kinetic energy $T = \frac{m_j \dot{\mathbf{r}}_j^2}{2}$.

- On to the second term $\text{II} \equiv m_j \dot{\mathbf{r}}_j \frac{d}{dt} \left[\frac{\partial \mathbf{r}_j}{\partial q_i} \right]$. First, we'll calculate the derivative $\frac{d}{dt} \frac{\partial \mathbf{r}_j}{\partial q_i}$, using the chain rule and remembering $\mathbf{r}_j = \mathbf{r}_j(q_1, \dots, q_n, t)$:

$$\frac{d}{dt} \frac{\partial \mathbf{r}_j}{\partial q_i} = \frac{\partial^2 \mathbf{r}_j}{\partial q_i \partial q_k} \dot{q}_k + \frac{\partial^2 \mathbf{r}_j}{\partial q_i \partial t} = \frac{\partial}{\partial q_i} \left[\frac{\partial \mathbf{r}_j}{\partial q_k} \dot{q}_k + \frac{\partial \mathbf{r}_j}{\partial t} \right] = \frac{\partial}{\partial q_i} \left[\frac{d\mathbf{r}_j}{dt} \right] = \frac{\partial \dot{\mathbf{r}}_j}{\partial q_i}$$

where we can factor out the derivative $\frac{\partial}{\partial q_i}$ in the second equality because $\frac{\partial \dot{q}_k}{\partial q_i} = 0$ (remember, the q_i and \dot{q}_i are different variables).

With the result $\frac{d}{dt} \frac{\partial \mathbf{r}_j}{\partial q_i} = \frac{\partial \dot{\mathbf{r}}_j}{\partial q_i}$ the term II, after a reverse engineered chain rule, becomes

$$\text{II} = m_j \dot{\mathbf{r}}_j \frac{d}{dt} \left[\frac{\partial \mathbf{r}_j}{\partial q_i} \right] = m_j \dot{\mathbf{r}}_j \frac{\partial \dot{\mathbf{r}}_j}{\partial q_i} = \frac{\partial}{\partial q_i} \left[\frac{m_j \dot{\mathbf{r}}_j^2}{2} \right] = \frac{\partial T}{\partial q_i}$$

- The hard work is done: we've eliminated the \mathbf{r}_j from both terms I and II in favor of q_i . Putting the pieces together, D'Alembert's principle in terms of q_i reads

$$\left[\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} - Q_j \right] \delta q_i = 0$$

Here's a summary of how we put the pieces back together, starting with $(\mathbf{F}_j^{(a)} - \dot{\mathbf{p}}_j) \cdot \delta \mathbf{r}_j = 0$

$$\begin{aligned} 0 &= \mathbf{F}_j^{(a)} \cdot \mathbf{r}_j - \dot{\mathbf{p}}_j \cdot \delta \mathbf{r}_j = Q_i \delta q_i - m_j \ddot{\mathbf{r}}_j \cdot \frac{\partial \mathbf{r}_j}{\partial q_i} \delta q_i = [Q_i - (\text{I} - \text{II})] \delta q_i \\ &= \left[Q_i - \frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) + \frac{\partial T}{\partial q_i} \right] \delta q_i \implies \left[\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} - Q_i \right] \delta q_i = 0 \end{aligned}$$

Now, here's why we transformed to generalized coordinates in the first place. Because the q_i independent of each other, the virtual displacements in separate coordinates δq_i and δq_j for $i \neq j$ are also independent. Because of the independence of the q_i , the equality in D'Alembert's principle holds for *all* $i = 1, \dots, n$ only if

$$\left[\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} - Q_i \right] = 0 \quad \text{or} \quad \frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} = Q_i$$

This is starting to look like the Lagrange's equations we saw in Section 2.1.

2.3.3 Deriving the Lagrange Equations: Take Two, Part Two

- The equations of motion

$$\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} - Q_j = 0$$

we just derived are in fact more general than the Lagrange equations $\frac{\partial L}{\partial q_i} = \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right)$ we saw in Section 2.1. To get to the more familiar form, we'll assume the forces acting on our system are conservative and thus can be derived from a scalar potential via

$$\mathbf{F} = -\nabla V(\mathbf{r}_1, \dots, \mathbf{r}_N) \quad \text{or, by components} \quad \mathbf{F}_j = -\frac{\partial V}{\partial \mathbf{r}_j}$$

In this case, the generalized forces $Q_i = \mathbf{F}_j \cdot \frac{\partial \mathbf{r}_j}{\partial \dot{q}_i}$ can be conveniently written

$$Q_i = -\frac{\partial V}{\partial \mathbf{r}_j} \cdot \frac{\partial \mathbf{r}_j}{\partial q_i} = -\frac{\partial V}{\partial q_i}$$

Plugging Q_i into the equations of motion then gives

$$\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial T}{\partial q_i} - \left(-\frac{\partial V}{\partial q_i} \right) = \frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial(T - V)}{\partial q_i} = 0$$

- One more assumption: we restrict our system to velocity independent potentials (which still covers most of the common interactions in nature, e.g. gravitational, electrostatic, etc...). A velocity-independent potential means $V \neq V(\dot{q}_i)$. The immediate implication $\frac{\partial V}{\partial \dot{q}_i} = 0$ lets us do a little trick, namely rewriting

$$\frac{\partial T}{\partial \dot{q}_j} = \frac{\partial(T - V)}{\partial \dot{q}_j}$$

The equations of motion then read

$$\frac{d}{dt} \left(\frac{\partial T}{\partial \dot{q}_i} \right) - \frac{\partial(T - V)}{\partial q_i} = \frac{d}{dt} \left[\frac{\partial(T - V)}{\partial \dot{q}_i} \right] - \frac{\partial(T - V)}{\partial q_i} = 0$$

Or, in terms of the Lagrangian $L \equiv T - V$, the familiar Lagrange equations

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} = 0$$

2.4 Noether's Theorem and Symmetries of the Lagrangian

2.4.1 Constants of Motion and Some Examples

- First, a definition: for a system with the generalized coordinates q_i , a function $F(q_i, \dot{q}_i, t)$ (i.e. a function of the generalized coordinates their time derivatives, and (possibly) time) is called a *constant of motion* or a *conserved quantity* if

$$\frac{dF}{dt} = \frac{\partial F}{\partial q_i} \dot{q}_i + \frac{\partial F}{\partial \dot{q}_i} \ddot{q}_i + \frac{\partial F}{\partial t} = 0$$

wherever $q_i(t)$ satisfy the Lagrange equations. In other words, its total time derivative must be zero for all allowed configurations of $q_i(t)$.

This definition tells us that any conserved quantity F is constant along a path followed by the system through its configuration space.

- **First Example:** If L doesn't depend explicitly on time (i.e. $\frac{\partial L}{\partial t} = 0$), the quantity

$$H = \dot{q}_i \frac{\partial L}{\partial \dot{q}_i} - L$$

is conserved. H is called the system's *Hamiltonian*. We'll hear all about it in Section 5. The Hamiltonian is associated with the system's total energy, and in fact when $\frac{\partial L}{\partial t} = 0$ we have $H = T + V$, which you can straightforwardly prove by recognizing $\dot{q}_i \frac{\partial L}{\partial \dot{q}_i} = 2T$ is twice the system's kinetic energy and substituting $L = T - V$.

To show the Hamiltonian is conserved when $L = L(q_i, \dot{q}_i)$ and $\frac{\partial L}{\partial t} = 0$ we calculate

$$\frac{dH}{dt} = \frac{d}{dt} \left[\dot{q}_i \frac{\partial L}{\partial \dot{q}_i} - L(q_i, \dot{q}_i) \right] = \ddot{q}_i \frac{\partial L}{\partial \dot{q}_i} + \dot{q}_i \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} \dot{q}_i - \frac{\partial L}{\partial \dot{q}_i} \ddot{q}_i$$

Because $\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) = \frac{\partial L}{\partial q_i}$ wherever q_i satisfy the Lagrange equations we have

$$\frac{dH}{dt} = \ddot{q}_i \frac{\partial L}{\partial \dot{q}_i} - \frac{\partial L}{\partial \dot{q}_i} \ddot{q}_i + \dot{q}_i \left(\frac{\partial L}{\partial q_i} \right) - \frac{\partial L}{\partial q_i} \dot{q}_i = 0$$

- **Second Example:** If $\frac{\partial L}{\partial q_i} = 0$ for some q_i , then the quantity

$$p_i \equiv \frac{\partial L}{\partial \dot{q}_i}$$

is a constant of motion. The proof is a straightforward application of the Lagrange equations $\frac{d}{dt} \frac{\partial L}{\partial \dot{q}_i} = \frac{\partial L}{\partial q_i}$ and the condition $\frac{\partial L}{\partial q_i} = 0$:

$$\frac{dp_i}{dt} \equiv \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) = \frac{\partial L}{\partial q_i} = 0$$

The quantity p_i is called the *generalized momentum* corresponding to the generalized coordinate q_i . In fact, we can re-write the Lagrange equations in the form

$$\dot{p}_i = \frac{\partial L}{\partial q_i}$$

The generalized momenta of Cartesian coordinates (e.g. x, y, z) agree with the Newtonian momentum $\mathbf{p} = m\mathbf{\dot{r}}$ but that is not the case for other types of generalized coordinates (e.g. angles or polar coordinates).

Generalized momenta play an important role in Hamiltonian dynamics, discussed in Section 5.

2.4.2 Symmetries of the Lagrangian

We'll need to define the concept of a symmetry before stating Noether's theorem.

- Consider a one-parameter family of maps of the form

$$q_i(t) \rightarrow Q_i(s, t) \quad \text{such that} \quad Q_i(0, t) = q_i(t)$$

parameterized by $s \in \mathbb{R}$. Such a map is called a *continuous symmetry of the Lagrangian* $L = L(q_i, \dot{q}_i, t)$ if

$$\frac{\partial}{\partial s} L(Q_i(s, t), \dot{Q}_i(s, t), t) = 0 \quad \text{for all } s \in \mathbb{R}$$

- That probably sounds abstract. Basically, we want to show that the Lagrangian for the coordinates q_i and Q_i is the same, i.e. $L(q_i, \dot{q}_i, t) = L(Q_i, \dot{Q}_i, t)$.

As an example we'll discuss symmetry of translation: the Lagrangian is the same under a translation of spatial coordinates. We write the map as

$$\mathbf{r}_i \rightarrow \mathbf{r}_i + s \hat{\mathbf{n}}$$

for some unit vector $\hat{\mathbf{n}}$ and scalar parameter s ; it is just a translation of the origin by $s \hat{\mathbf{n}}$. In the language of q_i and Q_i , we have

$$\begin{aligned} q_i &\leftrightarrow \mathbf{r}_i & \text{and} & & Q_i &\leftrightarrow \mathbf{r}_i + s \hat{\mathbf{n}} \\ \dot{q}_i &\leftrightarrow \dot{\mathbf{r}}_i & \text{and} & & \dot{Q}_i &\leftrightarrow \dot{\mathbf{r}}_i \end{aligned}$$

If the Lagrangian $L(q_i, \dot{q}_i, t)$ is

$$L(q_i, \dot{q}_i, t) = \frac{1}{2} \sum_i m_i \dot{\mathbf{r}}_i^2 - V(|\mathbf{r}_i - \mathbf{r}_j|)$$

then the Lagrangian $L(Q_i(s, t), \dot{Q}_i(s, t), t) \leftrightarrow L(\mathbf{r}_i + s \hat{\mathbf{n}}, \dot{\mathbf{r}}_i, t)$ would be

$$\begin{aligned} L(Q_i(s, t), \dot{Q}_i(s, t), t) &= \frac{1}{2} \sum_i m_i \dot{\mathbf{r}}_i^2 - V(|(\mathbf{r}_i + s \hat{\mathbf{n}}) - (\mathbf{r}_j + s \hat{\mathbf{n}})|) \\ &= \frac{1}{2} \sum_i m_i \dot{\mathbf{r}}_i^2 - V(|\mathbf{r}_i - \mathbf{r}_j|) \end{aligned}$$

It follows immediately that

$$\frac{\partial}{\partial s} L(Q_i(s, t), \dot{Q}_i(s, t), t) = 0 \quad \text{and} \quad L(Q_i(s, t), \dot{Q}_i(s, t), t) = L(q_i, \dot{q}_i, t)$$

The first equation is the formal definition that our space translation map $\mathbf{r}_i \rightarrow \mathbf{r}_i + s \hat{\mathbf{n}}$ is a symmetry of the Lagrangian, while the second equation gives a more intuitive reason why: the translated and untranslated Lagrangians are the same!

2.4.3 Statement of Noether's Theorem

- Here it is:

For every continuous symmetry of the Lagrangian, there exists a corresponding constant of motion $\left. \frac{\partial L}{\partial \dot{q}_i} \frac{\partial Q_i}{\partial s} \right|_{s=0}$

– **Proof:** Applying the chain rule to $L(Q_i(s, t), \dot{Q}_i(s, t), t)$ gives

$$\frac{\partial L}{\partial s} = \frac{\partial L}{\partial Q_i} \frac{\partial Q_i}{\partial s} + \frac{\partial L}{\partial \dot{Q}_i} \frac{\partial \dot{Q}_i}{\partial s}$$

By definition of a continuous symmetry $\frac{\partial L}{\partial s} = 0$ for all s . Choosing $s = 0$ and applying $Q_i|_{s=0} = q_i$ to $\frac{\partial L}{\partial s}$ we have:

$$0 \equiv \left. \frac{\partial L}{\partial s} \right|_{s=0} = \left. \frac{\partial L}{\partial Q_i} \frac{\partial Q_i}{\partial s} \right|_{s=0} + \left. \frac{\partial L}{\partial \dot{Q}_i} \frac{\partial \dot{Q}_i}{\partial s} \right|_{s=0}$$

- Next, we'll need the intermediate result $\left. \frac{\partial \dot{Q}_i}{\partial s} \right|_{s=0} = \left. \frac{d}{dt} \frac{\partial Q_i}{\partial s} \right|_{s=0}$. The result might already be intuitive; feel free to skip to the next bullet point. The key is the simplifications arising from $s = 0$. Differentiating $Q_i = Q_i(s, t)$ gives

$$\dot{Q}_i|_{s=0} = \left. \frac{dQ_i}{dt} \right|_{s=0} = \left. \frac{\partial Q_i}{\partial s} \dot{s} \right|_{s=0} + \left. \frac{\partial Q_i}{\partial t} \right|_{s=0} = 0 + \left. \frac{\partial Q_i}{\partial t} \right|_{s=0}$$

Differentiating \dot{Q}_i with respect to s then gives the desired result

$$\left. \frac{\partial \dot{Q}_i}{\partial s} \right|_{s=0} = \left. \frac{\partial^2 Q_i}{\partial s \partial t} \right|_{s=0} = \left. \frac{\partial}{\partial t} \frac{\partial Q_i}{\partial s} \right|_{s=0} = \left. \frac{d}{dt} \frac{\partial Q_i}{\partial s} \right|_{s=0}$$

where we could replace $\frac{\partial}{\partial t} \rightarrow \frac{d}{dt}$ in the last equality for the special case $s = 0$.

- Back to the main proof, where we left off with

$$0 \equiv \left. \frac{\partial L}{\partial s} \right|_{s=0} = \left. \frac{\partial L}{\partial q_i} \frac{\partial Q_i}{\partial s} \right|_{s=0} + \left. \frac{\partial L}{\partial \dot{q}_i} \frac{\partial \dot{Q}_i}{\partial s} \right|_{s=0}$$

Substituting $\frac{\partial L}{\partial q_i} = \frac{d}{dt} \frac{\partial L}{\partial \dot{q}_i}$ (the Lagrange equations), using $\left. \frac{\partial \dot{Q}_i}{\partial s} \right|_{s=0} = \left. \frac{d}{dt} \frac{\partial Q_i}{\partial s} \right|_{s=0}$ and recognizing a reverse product rule gives

$$0 \equiv \left. \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) \frac{\partial Q_i}{\partial s} \right|_{s=0} + \left. \frac{\partial L}{\partial \dot{q}_i} \frac{d}{dt} \left(\frac{\partial Q_i}{\partial s} \right) \right|_{s=0} = \left. \frac{d}{dt} \left[\frac{\partial L}{\partial \dot{q}_i} \frac{\partial Q_i}{\partial s} \right] \right|_{s=0}$$

The result $\left. \frac{d}{dt} \left[\frac{\partial L}{\partial \dot{q}_i} \frac{\partial Q_i}{\partial s} \right] \right|_{s=0} = 0$ proves that $\left. \frac{\partial L}{\partial \dot{q}_i} \frac{\partial Q_i}{\partial s} \right|_{s=0}$ evaluated at $s = 0$ is indeed a constant of motion.

2.4.4 Examples with Noether's Theorem

Linear Momentum and the Homogeneity of Space

Note that we'll briefly return to vector notation: it's easier for these examples.

- Consider an closed system of N particles with masses m_1, \dots, m_N and position vectors $\mathbf{r}_1, \dots, \mathbf{r}_N$. The system's Lagrangian is

$$L = \frac{1}{2} \sum_i m_i \dot{\mathbf{r}}_i^2 - V(|\mathbf{r}_i - \mathbf{r}_j|)$$

- We mentioned in Subsection 2.4.2 that the symmetry of space translation is written

$$\mathbf{r}_i \rightarrow \mathbf{r}_i + s \hat{\mathbf{n}}$$

which basically means shifting origin from $\mathbf{0}$ to $s \hat{\mathbf{n}}$. We showed in Subsection 2.4.2 that $\mathbf{r}_i \rightarrow \mathbf{r}_i + s \hat{\mathbf{n}}$ is a valid continuous symmetry of the Lagrangian L since

$$\frac{\partial}{\partial s} L(\mathbf{r}_i + s \hat{\mathbf{n}}, \dot{\mathbf{r}}_i, t) = 0 \quad \text{and} \quad L(\mathbf{r}_i, \dot{\mathbf{r}}_i, t) = L(\mathbf{r}_i + s \hat{\mathbf{n}}, \dot{\mathbf{r}}_i, t)$$

This is a formal way of saying that space is homogeneous and a translation of the system by $s \hat{\mathbf{n}}$ does nothing to the equations of motion.

- Applying Noether's theorem to the symmetry $\mathbf{r}_i \rightarrow \mathbf{r}_i + s \hat{\mathbf{n}}$ tells us that

$$\left. \frac{\partial L}{\partial \dot{q}_i} \frac{\partial Q_i}{\partial s} \right|_{s=0} = \sum_i^N \frac{\partial L}{\partial \dot{\mathbf{r}}_i} \cdot \hat{\mathbf{n}} = \sum_i^N \mathbf{p}_i \cdot \hat{\mathbf{n}}$$

is a constant of motion. The quantity $\sum_i \mathbf{p}_i \cdot \hat{\mathbf{n}}$ is the system's total linear momentum in the direction of the vector $\hat{\mathbf{n}}$. But since the choice of $\hat{\mathbf{n}}$ is arbitrary, $\sum_i \mathbf{p}_i \cdot \hat{\mathbf{n}}$ for all $\hat{\mathbf{n}}$, meaning the system's total linear momentum $\sum_i \mathbf{p}_i$ is conserved.

This is a profound result. Noether's theorem has shown us that homogeneity of space and the translational symmetry of the Lagrangian are responsible for conservation of linear momentum.

- Here is an intuitive and wonderfully informal explanation à la Tong: By homogeneity of space, one point in space is just as good as any other. Because one point is much the same as the next, why would a system of particles speed up (i.e. change it's linear momentum) to get "there" when "here" is just as good? This manifests itself as conservation of linear momentum.

Angular Momentum and the Isotropy of Space Isotropy of space is a fancy way of saying that space has the same properties in all directions. This manifests itself as invariance of space under rotations. But we'll have to be a bit more formal.

- Sticking with the closed system of N particles, if we rotate our system about some axis $\hat{\mathbf{n}}$, isotropy of space implies that all particles $\mathbf{r}_i \rightarrow \mathbf{r}'_i$ are rotated by the same amount. We can write the rotational mapping as an infinitesimal rotation

$$\mathbf{r}_i \rightarrow \mathbf{r}_i + \delta\theta \hat{\mathbf{n}} \times \mathbf{r}_i$$

parameterized by the infinitesimal angle $\delta\theta$.

Aside: We're working with infinitesimal rotations, which can be written in the form $\mathbf{r} \rightarrow \mathbf{r} + \delta\mathbf{r} = \mathbf{r} + \delta\theta \hat{\mathbf{n}} \times \mathbf{r}$ (instead of using a 3×3 rotation matrix). This makes the math easier. The justification is that we can reconstruct a finite rotation as a series of many infinitesimal rotations.

- If the Lagrangian is invariant under the rotational mapping, the conserved quantity is

$$\begin{aligned} \left. \frac{\partial L}{\partial \dot{q}_i} \frac{\partial}{\partial \delta\theta} [\mathbf{r}_i + \delta\theta \hat{\mathbf{n}} \times \mathbf{r}_i] \right|_{\delta\theta=0} &= \sum_i^N \frac{\partial L}{\partial \dot{\mathbf{r}}_i} \cdot (\hat{\mathbf{n}} \times \mathbf{r}_i) = \sum_i^N \mathbf{p}_i \cdot (\hat{\mathbf{n}} \times \mathbf{r}_i) \\ &= \sum_i^N \hat{\mathbf{n}} \cdot (\mathbf{r}_i \times \mathbf{p}_i) = \hat{\mathbf{n}} \cdot \sum_i^N \mathbf{L}_i = \hat{\mathbf{n}} \cdot \mathbf{L} \end{aligned}$$

which we recognize as the total linear momentum in the direction of $\hat{\mathbf{n}}$. The conclusion is the same as with linear momentum \mathbf{p} : since $\hat{\mathbf{n}} \cdot \mathbf{L}$ is conserved and $\hat{\mathbf{n}}$ is arbitrary, the system's total linear momentum \mathbf{L} is conserved.

Energy and the Homogeneity of Time

- Homogeneity of time is a statement that the laws of physics don't change with time; they hold just as well today as they did yesterday. We write the corresponding symmetry as

$$t \rightarrow t + s$$

which is just a shifting of time by the parameter s .

- If L is invariant under a time translation it cannot have explicit dependence on time, meaning

$$L = L(q_i, \dot{q}_i) \quad \text{and} \quad \frac{\partial L}{\partial t} = 0$$

But we have already shown in Subsection 2.4.1 that $\frac{\partial L}{\partial t} = 0$ implies the Hamiltonian

$$H = \frac{\partial L}{\partial \dot{q}_i} \dot{q}_i - L$$

is conserved. We also showed that $H = T + V$ is the system's total energy, so homogeneity of time leads us to conservation of energy.

- In fact, we have kind of cheated here. A derivation of energy conservation directly from Noether's theorem requires a more general statement of the theorem than given in this text. The better version is stated in terms of perturbations of the generalized coordinates and time. It's actually not too bad. I might add it later.

2.5 Small Oscillations and Stability

Let's say our system of N particles is resting in a state of equilibrium. Then we displace it ever so slightly. What happens? The answer lies in the study of small oscillations, which deals with small-displacement oscillations of a system of N particles about the equilibrium position. Happily, the dynamics is described by a simple and familiar system: the harmonic oscillator.

In fact, the motion of an N particle system near equilibrium can be modeled by a set of N *uncoupled* (i.e. independent) simple harmonic oscillators, each vibrating at a different frequency. More to come.

2.5.1 One Dimensional Example

- Let's start small, with a single particle and one degree of freedom q . In both Newtonian and Lagrangian mechanics, the motion is governed by a second-order differential equation with the general form

$$\ddot{q} = f(q)$$

for some (generally) non-linear function $f(q)$.

- At an equilibrium point $q = q_0$ the particle's acceleration \ddot{q} is zero, so the equation of motion must satisfy $f(q_0) = 0$ at q_0 .

If we start at rest at equilibrium with the initial conditions

$$q = q_0 \quad \text{and} \quad \dot{q} = 0$$

the system will remain in equilibrium for all time (formally, solving the equation of motion gives $q(t) = q_0$).

- But what happens if we start ever so slightly away from q_0 ? Allowing for time-dependence, we write

$$q(t) = q_0 + \eta(t)$$

where $\eta(t)$ are small displacements about equilibrium q_0 .

A first-order Taylor expansion of $\ddot{q} = f(q(t))$ gives

$$\ddot{q}(t) = f(q_0 + \eta(t)) \approx f(q_0) + f'(q_0)\eta(t) = f'(q_0)\eta(t)$$

where we've used the equilibrium requirement $f(q_0) = 0$. Since $\ddot{q} = \ddot{\eta}$, we have:

$$\ddot{\eta}(t) = f'(q_0)\eta(t)$$

- There are two possible outcomes for what happens next, depending on the sign of $f'(q_0)$.

1. If $f'(q_0) < 0$, the acceleration $\ddot{\eta}$ points towards q_0 , so q_0 is a position of stable equilibrium. The solution is sinusoidal oscillation about q_0 :

$$\eta(t) = \eta_0 \cos(\omega(t - t_0)) \quad \text{where} \quad \omega^2 = -f'(q_0)$$

The system undergoes stable harmonic oscillations about q_0 at frequency ω .

2. If $f'(q_0) > 0$, the acceleration $\ddot{\eta}$ points away from q_0 , so q_0 is a position of unstable equilibrium. The solution is exponential:

$$\eta(t) = Ae^{\lambda t} + Be^{-\lambda t} \quad \text{where} \quad \lambda = f'(q_0)$$

for general initial conditions (besides the very special case $A = 0$) η grows rapidly and the first-order Taylor approximation breaks down. The system is said to be *linearly unstable* at q_0 .

2.5.2 N Degrees of Freedom

- We now generalize to a system with N degrees of freedom. Our beginning steps are analogous to the one-dimensional case; the equations of motion are

$$\ddot{q}_i = f_i(q_1, \dots, q_N), \quad i = 1, \dots, n$$

where we've defined the vector of coordinates $\mathbf{q} = (q_1, \dots, q_N)$. An equilibrium position $\mathbf{q}^0 = (q_1^0, \dots, q_N^0)$ satisfies

$$\ddot{q}_i = 0 \quad \text{and} \quad f_i(\mathbf{q}^0) \equiv f_i(q_1^0, \dots, q_N^0) = 0 \quad \text{for all } i = 1, \dots, n$$

- Small displacements from the equilibrium point are written

$$q_i(t) = q_i^0 + \eta_i(t)$$

and a first-order Taylor expansion of $\ddot{\eta}_i = \ddot{q}_i$ gives

$$\begin{aligned} \ddot{\eta}_i(t) &= f_i(q_1^0 + \eta_1(t), \dots, q_N^0 + \eta_N(t)) \\ &\approx f_i(q_1^0, \dots, q_N^0) + \sum_{j=1}^N \frac{\partial}{\partial q_j} [f_i(q_1^0, \dots, q_N^0)] \eta_j \\ &= \sum_{j=1}^N \frac{\partial}{\partial q_j} [f_i(q_1^0, \dots, q_N^0)] \eta_j \quad (\text{since } f_i(q_1^0, \dots, q_N^0) \equiv 0) \end{aligned}$$

where the one-dimensional $f'(x_0)\eta$ is replaced with the total derivative $\frac{\partial f_i}{\partial q_j} \eta_j$ because $f_i = f_i(q_1, \dots, q_N)$ is now a multi-variable function.

- Since we're using multivariable calculus we'll switch to matrix form, which makes notation much cleaner. We introduce a vector of displacement $\boldsymbol{\eta} \in \mathbb{R}^N$ and the $(N \times N)$ matrix \mathbf{F} given by

$$\boldsymbol{\eta} = \begin{bmatrix} \eta_1 \\ \vdots \\ \eta_n \end{bmatrix} \quad \text{and} \quad \mathbf{F} = \left[\frac{\partial f_i}{\partial q_j} \right]_{q_i=q_i^0}$$

where \mathbf{F} is the matrix of partial derivatives *evaluated at equilibrium* $q_i = q_i^0$. The system's equation of motion near equilibrium then reads

$$\ddot{\boldsymbol{\eta}} = \mathbf{F}\boldsymbol{\eta}$$

- We have reduced the dynamics of the system to an eigenvalue problem. To finish our solution we must find \mathbf{F} 's eigenvectors and eigenvalues.

But first, a necessary bit of linear algebra theory:

- We cannot assume a priori that \mathbf{F} has a complete set of orthogonal eigenvectors with real eigenvalues. It will turn out, however, that any matrix \mathbf{F} derived from a physical Lagrangian system has all real eigenvalues. We'll prove this in a few paragraphs; for now we'll assume \mathbf{F} has real eigenvalues.
- In general, any matrix, including e.g. \mathbf{F} , has different right and left eigenvectors $\boldsymbol{\mu}_a$ and $\boldsymbol{\zeta}_a$ satisfying

$$\mathbf{F}\boldsymbol{\mu}_a = \lambda_a^2 \boldsymbol{\mu}_a \quad \boldsymbol{\zeta}_a^T \mathbf{F} = \lambda_a^2 \boldsymbol{\zeta}_a^T$$

where $a = 1, \dots, N$ and *there is no implied sum over a* . The left and right eigenvectors are orthonormal and satisfy $\boldsymbol{\mu}_a \cdot \boldsymbol{\zeta}_b = \delta_{ab}$.

Note that even though the eigenvectors differ, the eigenvalues λ_a^2 are the same for $a = 1, \dots, N$.

- Although the eigenvalues λ_a^2 are real for any physical system of interest, they are not always positive, allowing for the possibility of complex square roots $\sqrt{\lambda_a^2} = \pm \lambda_a \in \mathbb{C}$. That's okay; we'll see complex λ_a correspond to oscillation.
- Back to small oscillations, which we've reduced to the eigenvalue problem $\ddot{\boldsymbol{\eta}} = \mathbf{F}\boldsymbol{\eta}$. The most general solution to this equation is

$$\boldsymbol{\eta}(t) = \sum_a^N \boldsymbol{\mu}_a \left(A_a e^{\lambda_a t} + B_a e^{-\lambda_a t} \right)$$

where $\boldsymbol{\mu}_a$ are \mathbf{F} 's (left) eigenvectors, λ_a are \mathbf{F} 's eigenvalues, and A_a and B_a are integration constants determined by the initial conditions.

We have two possible behaviors for each a , depending on the sign on the eigenvalue λ_a^2

1. If $\lambda_a^2 < 0$ then $\pm \lambda_a \in \mathbb{C}$ is complex-valued and we write

$$\lambda_a = i\omega_a$$

for some real number $\omega_a \in \mathbb{R}$. This represents oscillation with the eigenfrequency ω_a about the equilibrium point q_a^N . The system is stable if displaced in the corresponding direction $\boldsymbol{\eta} = \boldsymbol{\mu}_a$.

2. If $\lambda_a^2 > 0$ then $\pm\lambda_a \in \mathbb{R}$ is real-valued (corresponding to an exponential solution $\eta_a(t) = A_a e^{\lambda_a t} + B_a e^{-\lambda_a t}$) and the system has a linear instability if displaced in the corresponding direction $\boldsymbol{\eta} = \boldsymbol{\mu}_a$.
- The eigenvectors $\boldsymbol{\mu}_a$ are called *normal modes* and the associated value ω_a is called the *eigenfrequency* or *normal frequency*. The equilibrium point $\mathbf{q}^0 = (q_1^0, \dots, q_N^0)$ is stable only if $\lambda_a^2 < 0$ for all $i = 1, \dots, n$, in which case the system oscillates about the equilibrium point as a linear superposition of the normal modes, each at the corresponding eigenfrequency.

In real form, the solution can be written

$$\boldsymbol{\eta}(t) = \sum_{a; \lambda_a^2 > 0} \boldsymbol{\mu}_a \left[A_a e^{\lambda_a t} + B_a e^{-\lambda_a t} \right] + \sum_{a; \lambda_a^2 < 0} A_a \boldsymbol{\mu}_a \cos(\omega_a(t - t_a))$$

where A_a, B_a and t_a are $2n$ integration constants determined by the initial conditions; the t_a replace the B_a for the oscillatory modes with $\lambda_a^2 < 0$.

Proof: \mathbf{F} 's Eigenvectors are Real

We want to show \mathbf{F} 's eigenvalues are real as long as \mathbf{F} is derived from a physical Lagrangian system (i.e. something we'd actually see in nature).

- Consider a general Lagrangian of the form

$$L = \frac{1}{2} \sum_{i,j} \mathbf{T}_{ij}(\mathbf{q}) \dot{q}_i \dot{q}_j - V(\mathbf{q})$$

where $\mathbf{T} : \mathbb{R}^N \rightarrow \mathbb{R}^{N \times N}$ is matrix encoding the coefficients of the kinetic energy and $\mathbf{q} \in \mathbb{R}^N$ is a vector of coordinates. We'll assume \mathbf{T} is invertible and positive definite for all \mathbf{q} . Finally, we'll need the potential energy matrix $\mathbf{V} : \mathbb{R}^N \rightarrow \mathbb{R}^{N \times N}$

$$\mathbf{V}(\mathbf{q}) = \begin{bmatrix} \frac{\partial^2 V}{\partial q_1 \partial q_1} & \cdots & \frac{\partial^2 V}{\partial q_1 \partial q_N} \\ \vdots & \ddots & \vdots \\ \frac{\partial^2 V}{\partial q_1 \partial q_N} & \cdots & \frac{\partial^2 V}{\partial q_N \partial q_N} \end{bmatrix}$$

Both \mathbf{T} and \mathbf{V} come out real-valued and symmetric by construction.

- Expanding about the equilibrium point \mathbf{q}^0 to linear order in displacement $\boldsymbol{\eta}$ gives

$$\mathbf{T}(\mathbf{q}^0) \ddot{\boldsymbol{\eta}} = -\mathbf{V}(\mathbf{q}^0) \boldsymbol{\eta} \quad \text{or} \quad \mathbf{T}_{ij}(\mathbf{q}^0) \ddot{\eta}_j = -\mathbf{V}_{ij}(\mathbf{q}^0) \eta_j \quad \text{for } i = 1, \dots, N$$

This might look frightening at first, but it is really just a set of equations of motion for the displacements $\ddot{\boldsymbol{\eta}}(t)$.

- We'll always be working with \mathbf{T} and \mathbf{V} evaluated at equilibrium \mathbf{q}^0 , so let's define the shorthand

$$\mathbf{T}^0 = \mathbf{T}(\mathbf{q}^0) \quad \text{and} \quad \mathbf{V}^0 = \mathbf{V}(\mathbf{q}^0)$$

Rearranging our result to get $\ddot{\boldsymbol{\eta}} = -(\mathbf{T}^0)^{-1} \mathbf{V}^0 \boldsymbol{\eta}$ and comparing to the equation of motion $\ddot{\boldsymbol{\eta}} = \mathbf{F} \boldsymbol{\eta}$, we see

$$\mathbf{F} = -(\mathbf{T}^0)^{-1} \mathbf{V}^0$$

- While \mathbf{T}^0 and \mathbf{V}^0 are both symmetric, they are not necessarily simultaneously diagonalizable (this occurs only if \mathbf{T}^0 is a scalar multiple of the identity matrix \mathbf{I}), which means that \mathbf{F} is not symmetric in general, as mentioned in the earlier interlude on linear algebra theory.

However, we can show \mathbf{F} has real eigenvalues. Since $\mathbf{F} = -(\mathbf{T}^0)^{-1}\mathbf{V}^0$ we have

$$\mathbf{F}\mathbf{u} = \lambda^2\boldsymbol{\mu} \implies \mathbf{V}^0\boldsymbol{\mu} = -\lambda^2\mathbf{T}^0\boldsymbol{\mu}$$

Multiplying the equation by complex conjugate $\bar{\boldsymbol{\mu}}$ gives

$$\bar{\boldsymbol{\mu}}\mathbf{V}^0\boldsymbol{\mu} = -\lambda^2\bar{\boldsymbol{\mu}}\mathbf{T}^0\boldsymbol{\mu}$$

Since \mathbf{V}^0 and \mathbf{T}^0 are both symmetric, the quantities $\bar{\boldsymbol{\mu}}\mathbf{V}^0\boldsymbol{\mu}$ and $\bar{\boldsymbol{\mu}}\mathbf{T}^0\boldsymbol{\mu}$ are both real, which suggests λ^2 must be real too.

In principle if both matrix quantities were zero the trivial equality $0 = 0$ would hold even for $\lambda^2 \in \mathbb{C}$. But this can't happen, since we've assumed \mathbf{T} is positive definite we know $\bar{\boldsymbol{\mu}}\mathbf{T}^0\boldsymbol{\mu} > 0$. It follows that \mathbf{F} 's eigenvalues λ^2 are indeed real.

2.5.3 Direct Lagrangian Mechanics Approach

- We consider a system of N degrees of freedom with the coordinates $\mathbf{q} = (q_1, \dots, q_N)$ whose Lagrangian can be written

$$L = T - V = \frac{1}{2} \sum_{i,j} T_{ij}(\mathbf{q}) \dot{q}_i \dot{q}_j - V$$

where \mathbf{T} is an $(N \times N)$ matrix encoding the coefficients of the kinetic energy. By construction \mathbf{T} turns out to be symmetric, so $T_{ij} = T_{ji}$. We'll assume $\frac{\partial L}{\partial t} = 0$, so the total energy is conserved and can be written

$$H = \frac{\partial L}{\partial \dot{q}_i} \dot{q}_i - L = T + V$$

- We assume there exists a stationary solution of Lagrange's equations at the equilibrium position

$$\mathbf{q}(t) = \mathbf{q}^0 \quad \text{or, by components} \quad q_i(t) = q_i^0 \quad \text{for } i = 1, \dots, N$$

At such a point the system is stationary so $\dot{q}_i = 0$ for all i and the kinetic energy term vanishes from the Lagrangian. We have

$$L|_{\mathbf{q}^0} = -V|_{\mathbf{q}^0} \quad \text{and} \quad \left. \frac{\partial L}{\partial q_i} \right|_{\mathbf{q}^0} = - \left. \frac{\partial V}{\partial q_i} \right|_{\mathbf{q}^0} = 0$$

- Stationary states \mathbf{q}^0 can be absolutely stable, unstable, or metastable.
 - *Absolutely stable states* satisfy

$$V(\mathbf{q}^0) < V(\mathbf{q}) \quad \text{for all } \mathbf{q} \neq \mathbf{q}^0$$

In other words, absolutely stable states are positions of absolute minimum of the potential energy V .

- *Metastable states* must satisfy the same condition, but only in a neighborhood of \mathbf{q}^0

$$V(\mathbf{q}^0) < V(\mathbf{q}) \quad \text{for } \mathbf{q} \text{ near } \mathbf{q}^0$$

In other words, metastable states are local minima of the potential energy V .

- *Unstable states* satisfy $V(\mathbf{q}^0) > V(\mathbf{q})$. Such states are maxima of the potential energy, and the force $\mathbf{F} = -\nabla V$ pushes particles away from the equilibrium position. Such states are not of interest in the study of small oscillations.
- We'll deal with metastable and absolutely stable states. We denote small displacements from equilibrium by $\boldsymbol{\eta}(t)$ and write

$$\mathbf{q}(t) = \mathbf{q}^0 + \boldsymbol{\eta}(t) \quad \text{or, by components} \quad q_i(t) = q_i^0(t) + \eta_i(t)$$

We then assume $V(\mathbf{q})$ is an analytic function of the coordinates \mathbf{q} , which is just a guarantee we can expand it into a Taylor series about equilibrium \mathbf{q}^0 . To second order in \mathbf{q} we have

$$\begin{aligned} V(\mathbf{q}) &\approx V(\mathbf{q}^0) + \left. \frac{\partial V}{\partial q_i} \right|_{\mathbf{q}^0} \eta_i + \frac{1}{2} \left. \frac{\partial^2 V}{\partial q_i \partial q_j} \right|_{\mathbf{q}^0} \eta_i \eta_j \\ &= V(\mathbf{q}^0) + \frac{1}{2} \left. \frac{\partial^2 V}{\partial q_i \partial q_j} \right|_{\mathbf{q}^0} \eta_i \eta_j \end{aligned}$$

where the linear term vanishes by the equilibrium condition $\left. \frac{\partial V}{\partial q_i} \right|_{\mathbf{q}^0} = 0$. The metastable equilibrium condition $V(\mathbf{q}) > V(\mathbf{q}^0)$ then implies

$$\frac{1}{2} \left. \frac{\partial^2 V}{\partial q_i \partial q_j} \right|_{\mathbf{q}^0} \eta_i \eta_j > 0 \iff \frac{1}{2} V_{ij} \eta_i \eta_j > 0$$

where the second expression is written in terms of the matrix \mathbf{V} of potential energy second derivatives

$$\mathbf{V} = \left[\begin{array}{ccc} \left. \frac{\partial^2 V}{\partial q_1 \partial q_1} \right|_{\mathbf{q}^0} & \cdots & \left. \frac{\partial^2 V}{\partial q_1 \partial q_N} \right|_{\mathbf{q}^0} \\ \vdots & \ddots & \vdots \\ \left. \frac{\partial^2 V}{\partial q_1 \partial q_N} \right|_{\mathbf{q}^0} & \cdots & \left. \frac{\partial^2 V}{\partial q_N \partial q_N} \right|_{\mathbf{q}^0} \end{array} \right]_{\mathbf{q}=\mathbf{q}^0}$$

By construction \mathbf{V} comes out to be symmetric and positive-definite, which will be important later on.

- We can then write the system's Lagrangian near equilibrium in the matrix notation

$$L = \frac{1}{2} \mathbf{T}_{ij} \dot{\eta}_i \dot{\eta}_j - \frac{1}{2} V_{ij} \eta_i \eta_j + V^0$$

where $\mathbf{T} = \mathbf{T}|_{\mathbf{q}^0}$. Lagrange's equations for the displacements η_i are then written

$$\mathbf{T}_{ij} \ddot{\eta}_j + V_{ij} \eta_j = 0 \quad \text{or} \quad \mathbf{T} \ddot{\boldsymbol{\eta}} = \mathbf{V} \boldsymbol{\eta}$$

- The basic solution of the equations of motion $\mathbf{T}_{ij} \ddot{\eta}_j + V_{ij} \eta_j = 0$ are so called *characteristic oscillations*, which are harmonic oscillations of all the displacement components $\eta_i(t)$ with the same frequency ω , represented by the ansatz

$$\eta_j(t) = \alpha a_j e^{-i\omega t}$$

where α describes the amplitude of the oscillations and the unit vector $\mathbf{a} = (a_1, \dots, a_N)$ encodes the direction of each displacement component η_i .

To get a relationship between ω^2 and \mathbf{a} , we calculate $\ddot{\eta}_j = -\omega^2 \alpha a_j e^{-\omega t}$, plug this in to the equation of motion, cancel like terms and rearrange to get

$$V_{ij}a_j = \omega^2 T_{ij}a_j \quad \text{or} \quad \mathbf{V}\mathbf{a} = \omega^2 \mathbf{T}\mathbf{a}$$

- The result $\mathbf{V}\mathbf{a} = \omega^2 \mathbf{T}\mathbf{a}$ is important: we have reduced the dynamics of small oscillations to a generalized eigenvalue problem involving the matrices \mathbf{T} and \mathbf{V} . Because the matrices are symmetric, the eigenvalues ω^2 are all real and because \mathbf{T} is positive definite, we know the eigenvalues will be positive $\omega^2 > 0$.

Note, however, that this is in general *not* a simple eigenvalue problem of the form $\mathbf{A}\mathbf{x} = \omega^2 \mathbf{x}$; there are two matrices involved!

2.5.4 The Simplified Eigenvalue Problem

- Let's start on more familiar territory with the simple eigenvalue problem of the general form $\mathbf{A}\mathbf{x} = \omega^2 \mathbf{x}$. This occurs if $\mathbf{T} = T\mathbf{I}$, i.e. if \mathbf{T} is diagonal with identical scalar elements T . In this case, the problem reduces to the familiar eigenvalue problem $\mathbf{V}\mathbf{a} = \omega^2 T\mathbf{a}$.
- Because \mathbf{V} is symmetric (and the \mathbf{a} are normalized by construction) the eigenvectors are orthogonal:

$$\mathbf{a}_i^T \cdot \mathbf{a}_j = \delta_{ij}$$

We write the eigenvectors \mathbf{a}_i in the $N \times N$ matrix \mathbf{A}

$$\mathbf{A} = [\mathbf{a}_1, \dots, \mathbf{a}_N] = \begin{bmatrix} a_{11} & \cdots & a_{1N} \\ \vdots & \ddots & \vdots \\ a_{N1} & \cdots & a_{NN} \end{bmatrix}$$

Because the \mathbf{a}_i are orthonormal (i.e. $\mathbf{a}_i \cdot \mathbf{a}_j = \delta_{ij}$) we have

$$\mathbf{A}^T \mathbf{A} = \mathbf{I} \implies \mathbf{A} = \mathbf{A}^{-1}$$

- Some new notation coming fast: for \mathbf{V} 's eigenvalues we'll write $\omega^2 \equiv \lambda$, so the eigenvalue problem reads

$$\mathbf{V}\mathbf{a} = \lambda \mathbf{T}\mathbf{a} = \lambda T\mathbf{I} \quad \text{or} \quad (\mathbf{V} - \lambda T\mathbf{I})\mathbf{a} = 0$$

(We're still assuming $\mathbf{T} = T\mathbf{I}$). \mathbf{A} 's eigenvalues are $\lambda_i T$ for $i = 1, \dots, N$, which we'll write $\tilde{\lambda}_i \equiv \lambda_i T$ for shorthand, so our problem now reads $(\mathbf{V} - \tilde{\lambda}\mathbf{I})\mathbf{a} = 0$. Since we require $\mathbf{a} \neq \mathbf{0}$, the equation is satisfied if $|\mathbf{V} - \tilde{\lambda}\mathbf{I}| = 0$.

- In terms of left and right eigenvectors, the eigenvalue problem is

$$(\mathbf{V} - \tilde{\lambda}_i \mathbf{I})\mathbf{a}_i = 0 \quad \text{and} \quad \mathbf{a}_j^T (\mathbf{V} - \tilde{\lambda}_j \mathbf{I}) = 0$$

We then combine these two equations into the matrix equation

$$\mathbf{A}^T \mathbf{V} \mathbf{A} - \mathbf{A}^T \mathbf{A} \tilde{\Lambda} = 0$$

where $\mathbf{\Lambda}$ is a diagonal matrix whose diagonal elements are the eigenvalues $\tilde{\lambda}_i$. Because the eigenvectors \mathbf{a}_i are orthonormal and $\mathbf{A}^T \mathbf{A} = \mathbf{I}$, the matrix $\mathbf{\Lambda}$ is also unitary.

From the matrix equation $\mathbf{A}^T \mathbf{V} \mathbf{A} = \mathbf{A}^T \mathbf{A} \mathbf{\Lambda}$, we can interpret the eigenvector matrix \mathbf{A} as a transformation matrix that transforms \mathbf{V} into the diagonal form $\mathbf{\Lambda}$, i.e.

$$\mathbf{A}^T \mathbf{V} \mathbf{A} = \mathbf{A}^{-1} \mathbf{V} \mathbf{A} = \mathbf{A}^{-1} \mathbf{A} \mathbf{\Lambda} = \mathbf{\Lambda}$$

where we've used $\mathbf{V} \mathbf{A} = \mathbf{A} \mathbf{\Lambda}$ and the orthogonality property $\mathbf{A}^T = \mathbf{A}^{-1}$.

2.5.5 Generalized Eigenvalue Problem

- On to the generalized eigenvalue problem with $\mathbf{T} \neq T\mathbf{I}$

$$\mathbf{V} \mathbf{a} = \omega^2 \mathbf{T} \mathbf{a} = \lambda \mathbf{T} \mathbf{a}$$

where we'll continue with the notation $\lambda \equiv \omega^2$ for the eigenvalues. Both \mathbf{V} and \mathbf{T} are real and symmetric, so the eigenvalues λ and eigenvectors \mathbf{a} are real.

- As before, we write $(\mathbf{V} - \lambda \mathbf{T}) \mathbf{a} = 0$ and require $|\mathbf{V} - \lambda \mathbf{T}| = 0$, which disregards the trivial solution $\mathbf{a} = 0$.

In terms of left and right eigenvectors, the eigenvalue problem reads

$$\mathbf{V} \mathbf{a}_i = \lambda_i \mathbf{T} \mathbf{a}_i \quad \text{and} \quad \mathbf{a}_j^T \mathbf{V} = \lambda_j \mathbf{a}_j^T \mathbf{T}$$

Next, we'll prove the eigenvectors satisfy a generalized orthogonality. Multiplying the first equation from the left by \mathbf{a}_j^T and the second equation from the right by \mathbf{a}_i gives

$$\mathbf{a}_j^T \mathbf{V} \mathbf{a}_i = \lambda_i \mathbf{a}_j^T \mathbf{T} \mathbf{a}_i \quad \text{and} \quad \mathbf{a}_j^T \mathbf{V} \mathbf{a}_i = \lambda_j \mathbf{a}_j^T \mathbf{T} \mathbf{a}_i$$

Subtracting the equations leads to $(\lambda_i - \lambda_j) \mathbf{a}_j^T \mathbf{T} \mathbf{a}_i = 0$, which implies

$$\mathbf{a}_j^T \mathbf{T} \mathbf{a}_i = 0 \quad \text{for} \quad \lambda_i \neq \lambda_j$$

Even for degenerate eigenvalues $\lambda_i = \lambda_j, i \neq j$, we can choose the eigenvectors so they satisfy the orthogonality relation. More so, because \mathbf{T} is positive definite, its diagonal elements must be positive and nonzero. This allows us to choose the eigenvectors to satisfy the generalized normalization condition

$$\mathbf{a}_i^T \mathbf{T} \mathbf{a}_i = 1$$

- Together, the generalized orthogonality ($\mathbf{a}_j^T \mathbf{T} \mathbf{a}_i = 0$) and normalization ($\mathbf{a}_i^T \mathbf{T} \mathbf{a}_i = 1$) conditions mean that the eigenvector matrix \mathbf{A} transforms the kinetic energy matrix \mathbf{T} into the identity matrix

$$\mathbf{A}^T \mathbf{T} \mathbf{A} = \mathbf{I}$$

Like for the simpler case, we can now write the eigenvalue problem as a matrix equation

$$\mathbf{V} \mathbf{A} = \mathbf{T} \mathbf{A} \mathbf{\Lambda}$$

where $\mathbf{\Lambda}$ is a diagonal matrix whose elements are the eigenvalues λ_i .

As before, we can view \mathbf{A} as a transformation matrix that transforms \mathbf{V} into the diagonal form $\mathbf{\Lambda}$. Using $\mathbf{V} \mathbf{A} = \mathbf{T} \mathbf{A} \mathbf{\Lambda}$ and $\mathbf{A}^T \mathbf{T} \mathbf{A} = \mathbf{I}$ we have

$$\mathbf{A}^T \mathbf{V} \mathbf{A} = \mathbf{A}^T \mathbf{T} \mathbf{A} \mathbf{\Lambda} = \mathbf{\Lambda}$$

In the general case, however, \mathbf{A} has the additional property of transforming \mathbf{T} into the identity matrix \mathbf{I} .

2.5.6 Normal Coordinates

- A few sections ago we derived the equations of motion governing small oscillations

$$\mathbf{T}_{ij}\ddot{\eta}_j + \mathbf{V}_{ij}\eta_j = 0$$

and said the basic solutions were the characteristic oscillations

$$\eta_j(t) = \alpha a_j e^{-i\omega t}$$

The general solution is a linear superposition of the characteristic oscillations.

- The general solution is best written in terms of *normal coordinates*, which replace the small displacements η_i of the original coordinates q_i with the amplitudes α_j of the characteristic oscillations \mathbf{a}_j . We write the transformation to the normal coordinates as

$$\eta_i(t) = \alpha_j(t) a_{ji} = a_{ji} \alpha_j(t)$$

where a_{ji} is the i -th component of the j -th eigenvector \mathbf{a}_j . Note also that a_{ji} corresponds to the element A_{ij} of the eigenvector matrix $\mathbf{A} = [\mathbf{a}_1, \dots, \mathbf{a}_N]$.

We can thus write the transformation to normal coordinates in the matrix form

$$\boldsymbol{\eta}(t) = \mathbf{A}\boldsymbol{\alpha}(t) \quad \text{or} \quad \boldsymbol{\eta}^T(t) = \boldsymbol{\alpha}^T(t)\mathbf{A}^T$$

- The normal coordinates allow us to concisely write the system's Lagrangian

$$L = \frac{1}{2} (\mathbf{T}_{ij}\dot{\eta}_i\dot{\eta}_j - \mathbf{V}_{ij}\eta_i\eta_j) + V^0 = \frac{1}{2} (\dot{\boldsymbol{\eta}}^T \mathbf{T} \dot{\boldsymbol{\eta}} - \boldsymbol{\eta}^T \mathbf{V} \boldsymbol{\eta}) + V^0$$

in terms of the eigenvalues ω_j^2 . First, let's redefine the potential so $V^0 = 0$. Then, using $\boldsymbol{\eta} = \mathbf{A}\boldsymbol{\alpha}$ and the identities $\mathbf{A}^T \mathbf{T} \mathbf{A} = \mathbf{I}$ and $\mathbf{A}^T \mathbf{V} \mathbf{A} = \boldsymbol{\Lambda} = \text{diag}(\omega_j^2)$ we have

$$\begin{aligned} L &= \frac{1}{2} (\dot{\boldsymbol{\eta}}^T \mathbf{T} \dot{\boldsymbol{\eta}} - \boldsymbol{\eta}^T \mathbf{V} \boldsymbol{\eta}) = \frac{1}{2} (\dot{\boldsymbol{\alpha}}^T \mathbf{A}^T \mathbf{T} \mathbf{A} \dot{\boldsymbol{\alpha}} - \boldsymbol{\alpha}^T \mathbf{A}^T \mathbf{V} \mathbf{A} \boldsymbol{\alpha}) \\ &= \frac{1}{2} (\dot{\boldsymbol{\alpha}}^T \mathbf{I} \dot{\boldsymbol{\alpha}} - \boldsymbol{\alpha}^T \boldsymbol{\Lambda} \boldsymbol{\alpha}) = \frac{1}{2} (\dot{\boldsymbol{\alpha}}^T \dot{\boldsymbol{\alpha}} - \boldsymbol{\alpha}^T \boldsymbol{\Lambda} \boldsymbol{\alpha}) \\ &= \frac{1}{2} \sum_j (\dot{\alpha}_j^2 - \omega_j^2 \alpha_j^2) \end{aligned}$$

- The time dependence of each normal coordinate α_j is the harmonic oscillation

$$\alpha_j(t) = \alpha_{j0} \cos(\omega_j t - \delta)$$

where the amplitude α_{j0} and phase δ are integration constants determined by the initial conditions.

The Lagrangian is thus a sum of $j = 1, \dots, N$ uncoupled harmonic oscillators $\alpha_j(t)$, each vibrating at the frequency ω_j , where the amplitude of oscillation is the j -th normal coordinate α_j .

2.6 A Concluding Example: A Particle in an Electromagnetic Field

- This example is important because of the velocity-dependent nature of the electromagnetic interaction (i.e. the Lorentz force). So far, we've worked almost exclusively with velocity-independent forces in our discussion of Lagrangian mechanics. Now we'll see how the Lagrangian formalism handles a velocity-dependent force.
- We'll return to vector coordinates for this discussion. The electric field \mathbf{E} and magnetic field \mathbf{B} can be written in terms of a scalar potential $\phi(\mathbf{r}, t)$ and a vector potential $\mathbf{A}(\mathbf{r}, t)$, respectively

$$\mathbf{E} = -\nabla\phi - \frac{\partial\mathbf{A}}{\partial t} \quad \text{and} \quad \mathbf{B} = \nabla \times \mathbf{A}$$

The Lagrangian for a particle of charge e in an electromagnetic field is (currently without derivation)

$$L = \frac{1}{2}m\dot{\mathbf{r}}^2 - e(\phi - \dot{\mathbf{r}} \cdot \mathbf{A})$$

Note the Lagrangian splits into the usual kinetic energy term $\frac{1}{2}m\dot{\mathbf{r}}^2$ and an additional term due to the EM interaction. The momentum \mathbf{p} conjugate to the coordinate \mathbf{r} is

$$\mathbf{p} = \frac{\partial L}{\partial \dot{\mathbf{r}}} = m\dot{\mathbf{r}} + e\mathbf{A}$$

Note that the momentum is not simply the Newtonian $m\dot{\mathbf{r}}$, but is modified in the presence of an EM field.

- With Lagrangian in hand, we can calculate Lagrange's equations

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{\mathbf{r}}} \right) - \frac{\partial L}{\partial \mathbf{r}} = \frac{d}{dt} (m\dot{\mathbf{r}} + e\mathbf{A}) - e(\nabla\phi - \nabla(\dot{\mathbf{r}} \cdot \mathbf{A}))$$

In terms of components and the indices $a, b \in \{1, 2, 3\}$ of the Cartesian components the equation of motion reads

$$m\ddot{r}^a = -e \left(\frac{\partial\phi}{\partial r^a} + \frac{\partial A_a}{\partial t} \right) + e \left(\frac{\partial A_b}{\partial r^a} - \frac{\partial A_a}{\partial r^b} \right) \dot{r}^b$$

- In terms of indices and components, the definitions of the \mathbf{E} and \mathbf{B} fields are

$$E_a = -\frac{\partial\phi}{\partial r^a} - \frac{\partial A_a}{\partial t} \quad \text{and} \quad B_a = \epsilon_{cab} \frac{\partial A_b}{\partial r^a} \iff \epsilon_{abc} B_c = \frac{\partial A_b}{\partial r^a} - \frac{\partial A_a}{\partial r^b}$$

But these expressions agree exactly with the terms in parentheses in the equation of motion! Substitution, the equation of motion becomes

$$m\ddot{r}^a = eE_a + \epsilon_{abc} e \dot{r}^b B_c \quad \text{or} \quad m\ddot{\mathbf{r}} = e(\mathbf{E} + \dot{\mathbf{r}} \times \mathbf{B})$$

which is the Lorentz force law!

3 Dynamics of Central Force Motion

The chapter is dedicated to the motion of particles in a central force potential in three spatial dimensions. We will begin with Newtonian mechanics and use the Lagrangian formalism later in the chapter.

3.1 The Central Force Problem

3.1.1 Two Body Problem

The two body problem is the name given to solving the dynamics of two mutually interacting particles. It turns out we can reduce the two body problem into an equivalent one-body problem by using conserved quantities. Here's how it works:

- First, let's define the actors involved. We consider two bodies of mass m_1 and m_2 whose position is given by the position vectors \mathbf{r}_1 and \mathbf{r}_2 . We denote the total mass by $M = m_1 + m_2$ and define the relative separation as

$$\mathbf{r} = \mathbf{r}_1 - \mathbf{r}_2$$

Finally, we define the center of mass \mathbf{R} with the equation

$$M\mathbf{R} = m_1\mathbf{r}_1 + m_2\mathbf{r}_2$$

- We begin simplifying the problem by writing

$$\mathbf{r}_1 = \mathbf{R} + \frac{m_2}{M}\mathbf{r} \quad \text{and} \quad \mathbf{r}_2 = \mathbf{R} - \frac{m_1}{M}\mathbf{r}$$

Under the assumption that the two bodies are isolated, there is no external force on the system, which means the center of mass travels with constant velocity, i.e. $\ddot{\mathbf{R}} = 0$.

- Meanwhile, the relative motion of the two particles is governed by

$$\ddot{\mathbf{r}} = \ddot{\mathbf{r}}_1 - \ddot{\mathbf{r}}_2 = \frac{\mathbf{F}_{12}}{m_1} - \frac{\mathbf{F}_{21}}{m_2} = \frac{m_1 + m_2}{m_1 m_2}$$

where \mathbf{F}_{12} and \mathbf{F}_{21} are the forces exerted by the second body on the first body and vice versa, respectively. The last equality uses Newton's third law, which tells us $\mathbf{F}_{12} = -\mathbf{F}_{21}$.

- It is customary (and efficient) to write the equation of motion for $\ddot{\mathbf{r}}$ in the condensed form

$$\mu\ddot{\mathbf{r}} = \mathbf{F}_{12}$$

where we have defined the *reduced mass*

$$\mu = \frac{m_1 m_2}{m_1 + m_2}$$

Keep in mind that when one of the particles is much more massive than the other, the reduced mass nearly equals the mass of the lighter particle (e.g. $m_2 \gg m_1 \implies \mu \approx m_1$). Such situation come up often in practice (e.g. a satellite orbit an planet or an electron orbiting a hydrogen atom).

- The equation $\mu = \frac{m_1 m_2}{m_1 + m_2}$ is a great result: we have encoded the dynamics of our two body system in a single vector \mathbf{r} , effectively turning our two body problem into a one body problem (we could do this because the velocity of the center of mass is conserved).

To proceed further, we have to make some assumptions about the nature of the interaction \mathbf{F}_{12} . If \mathbf{F}_{12} is a *central force*, meaning it depends only on the magnitude $r = |\mathbf{r}|$ of the relative separation, we can make a great deal more progress in analyzing the two body problem. We will dedicate the rest of this section to analyzing central forces.

3.1.2 Basics of the Central Force Problem

- We start with a single particle in a central force potential (a potential that depends only on the distance $r = |\mathbf{r}|$ from the origin.) The equation of motion for our one-particle system is

$$m\ddot{\mathbf{r}} = -\nabla V(r)$$

where $V = V(r)$ is a central potential. Both the gravitational and electrostatic interactions have central potentials, so the results of this section have wide applicability in classical physics.

- The central problem is solved using angular momentum

$$\mathbf{L} = \mathbf{r} \times \mathbf{p} = m\mathbf{r} \times \dot{\mathbf{r}}$$

Angular momentum is conserved in a central potential. The proof is straightforward; we just have to show $\frac{d\mathbf{L}}{dt} = 0$

$$\frac{d\mathbf{L}}{dt} = m\dot{\mathbf{r}} \times \dot{\mathbf{r}} + m\mathbf{r} \times \ddot{\mathbf{r}} = 0 - \mathbf{r} \times \nabla V(r) = 0$$

where the last equality holds because \mathbf{r} and $\nabla V(r)$ are parallel for a central potential.

- The conservation of angular momentum has far-reaching consequences for the central force problem. For one, it means all motion takes place in a plane. How to show this? We know \mathbf{L} is conserved, so its direction and magnitude are fixed and unchanging. By definition $\mathbf{L} = \mathbf{r} \times \mathbf{p}$ we know \mathbf{L} and position \mathbf{r} are perpendicular, so

$$\mathbf{L} \cdot \mathbf{r} = 0$$

and the position of the particle is always perpendicular to \mathbf{L} . But because \mathbf{L} is *fixed* and \mathbf{r} is always perpendicular to \mathbf{L} , \mathbf{r} must lie in a *plane* perpendicular to \mathbf{L} . By the same argument, starting with $\mathbf{L} = m\mathbf{r} \times \dot{\mathbf{r}}$, we have $\mathbf{L} \cdot \dot{\mathbf{r}} = 0$, so the particle's velocity also lies in the same plane perpendicular to \mathbf{L} .

For central forces, three dimensional dynamics is reduced to dynamics in a plane because of the conservation of \mathbf{L} .

3.1.3 Polar Coordinates and the Central Force Problem

- Because motion under central forces lies in a plane, the dynamics is best analyzed with polar coordinates. This section explains how we will set up our polar coordinate system.

- The system's conserved angular momentum \mathbf{L} provides a natural reference point, so we rotate our coordinate system so the angular momentum points in the positive z direction. The motion, which we have shown, lies in the plane perpendicular to \mathbf{L} , the occurs in our (x, y) plane. We define the usual polar coordinates

$$x = r \cos \theta \quad \text{and} \quad y = r \sin \theta$$

- The next step is to express a particle's velocity and acceleration in polar coordinates. To do this, we introduce the two unit vectors $\hat{\mathbf{r}}$

$\hat{\boldsymbol{\theta}}$;

they point in direction of increasing r and θ , respectively.

- This is best seen in a figure. In words, you can visualize $\hat{\mathbf{r}}$ pointing from the particle's position away from the origin, parallel to the position vector \mathbf{r} . $\hat{\boldsymbol{\theta}}$ is perpendicular to $\hat{\mathbf{r}}$ and points from the particle's position in the direction of increasing θ in a counterclockwise sense.

In Cartesian form, the unit vectors can be written in terms of θ as

$$\hat{\mathbf{r}} = \begin{bmatrix} \cos \theta \\ \sin \theta \end{bmatrix} \quad \text{and} \quad \hat{\boldsymbol{\theta}} = \begin{bmatrix} -\sin \theta \\ \cos \theta \end{bmatrix}$$

- Notice that the basis $\{\hat{\mathbf{r}}, \hat{\boldsymbol{\theta}}\}$ depends on the particle's coordinate θ . The change with respect to θ is

$$\frac{d\hat{\mathbf{r}}}{d\theta} = \begin{bmatrix} -\sin \theta \\ \cos \theta \end{bmatrix} = \hat{\boldsymbol{\theta}} \quad \text{and} \quad \frac{d\hat{\boldsymbol{\theta}}}{d\theta} = -\begin{bmatrix} \cos \theta \\ \sin \theta \end{bmatrix} = -\hat{\mathbf{r}}$$

- The derivatives tell us that if the particle moves in such a way that θ changes with time then the basis vectors $\{\hat{\mathbf{r}}, \hat{\boldsymbol{\theta}}\}$ also change with time. What does this mean for \mathbf{r} , $\dot{\mathbf{r}}$ and $\ddot{\mathbf{r}}$? By construction of $\hat{\mathbf{r}}$, the position vector is simply

$$\mathbf{r} = r \hat{\mathbf{r}}$$

The velocity is found with the product rule for derivatives

$$\dot{\mathbf{r}} = \dot{r} \hat{\mathbf{r}} + r \frac{d\hat{\mathbf{r}}}{dt}$$

This turns into a more useful expression with the chain rule $\frac{d\hat{\mathbf{r}}}{dt} = \frac{d\hat{\mathbf{r}}}{d\theta} \frac{d\theta}{dt}$ and identity $\frac{d\hat{\mathbf{r}}}{d\theta} = \hat{\boldsymbol{\theta}}$. We get

$$\dot{\mathbf{r}} = \dot{r} \hat{\mathbf{r}} + r \dot{\theta} \hat{\boldsymbol{\theta}}$$

The second term $r \dot{\theta} \hat{\boldsymbol{\theta}}$ comes from the rotation $\dot{\theta}$ of the basis vectors and is a consequence of a non-inertial frame of reference. Notice that for a stationary system $\dot{\theta} = 0$ recovers the familiar expression $\dot{\mathbf{r}} = \dot{r} \hat{\mathbf{r}}$.

- Next, acceleration. A time differentiation of $\dot{\mathbf{r}}$ gives

$$\ddot{\mathbf{r}} = \ddot{r} \hat{\mathbf{r}} + \dot{r} \frac{d\hat{\mathbf{r}}}{dt} + \dot{r} \dot{\theta} \hat{\boldsymbol{\theta}} + r \ddot{\theta} \hat{\boldsymbol{\theta}} + r \dot{\theta} \frac{d\hat{\boldsymbol{\theta}}}{dt}$$

Again, the chain rule and basic vector derivatives come in handy. Using the identities $\frac{d\hat{\mathbf{r}}}{dt} = \frac{d\hat{\mathbf{r}}}{d\theta} \frac{d\theta}{dt} = \dot{\theta} \hat{\boldsymbol{\theta}}$ and $\frac{d\hat{\boldsymbol{\theta}}}{dt} = \frac{d\hat{\boldsymbol{\theta}}}{d\theta} \frac{d\theta}{dt} = -\dot{\theta} \hat{\mathbf{r}}$ we get

$$\ddot{\mathbf{r}} = (\ddot{r} - r \dot{\theta}^2) \hat{\mathbf{r}} + (r \ddot{\theta} + 2\dot{r} \dot{\theta}) \hat{\boldsymbol{\theta}}$$

- As a quick example, consider uniform circular motion in the plane at radius r and constant angular speed $\dot{\theta} = \omega$. In this case (because $\dot{r} = 0$) velocity is

$$\dot{\mathbf{r}} = r\omega \hat{\theta} \quad v = |\dot{\mathbf{r}}| = r\omega$$

At uniform angular speed ($\dot{\omega} = \ddot{\theta} = 0$) the acceleration is

$$\ddot{\mathbf{r}} = -r\omega^2 \hat{\mathbf{r}} \quad a = |\ddot{\mathbf{r}}| = r\omega^2$$

which is just the familiar centripetal acceleration $a = \frac{v^2}{r}$.

3.1.4 More On Central Forces

- Recall the equation of motion $m\ddot{\mathbf{r}} = -\nabla V(r)$. Because $V = V(r)$ depends only on distance r from the origin, the gradient is

$$\nabla V = \frac{dV}{dr} \hat{\mathbf{r}}$$

In terms of the acceleration $\ddot{\mathbf{r}} = (\ddot{r} - r\dot{\theta}^2) \hat{\mathbf{r}} + (r\ddot{\theta} + 2\dot{r}\dot{\theta}) \hat{\theta}$, the central force equation of motion becomes

$$m(\ddot{r} - r\dot{\theta}^2) \hat{\mathbf{r}} + m(r\ddot{\theta} + 2\dot{r}\dot{\theta}) \hat{\theta} = -\frac{dV}{dr} \hat{\mathbf{r}}$$

- Notice the acceleration is resolved into a radial component $(\ddot{r} - r\dot{\theta}^2) \hat{\mathbf{r}}$ and tangential component $(r\ddot{\theta} + 2\dot{r}\dot{\theta}) \hat{\theta}$.

Because there is no $\hat{\theta}$ term on the right hand side of the equation (or, if you will, a $0 \cdot \hat{\theta}$ term), we have $(r\ddot{\theta} + 2\dot{r}\dot{\theta}) = 0$. The $\hat{\theta}$ component being zero should make sense, since central forces act only along the radial direction.

There is something cool just below the surface here. We just need to re-write $r\ddot{\theta} + 2\dot{r}\dot{\theta}$ as a time derivative.

$$0 \equiv (r\ddot{\theta} + 2\dot{r}\dot{\theta}) = \frac{1}{r} \frac{d}{dt} \implies \frac{d}{dt} = 0$$

In other words, the quantity $r^2\dot{\theta}$ is conserved in central force motion.

- But what does the conserved quantity $r^2\dot{\theta}$ actually represent? In fact, we've met it already: it is angular momentum in disguise. More specifically, it corresponds to the *magnitude* of angular momentum.

To reveal this, we first rewrite \mathbf{L} in terms of the basis vectors $\{\hat{\mathbf{r}}, \hat{\theta}\}$:

$$\mathbf{L} = m\mathbf{r} \times \dot{\mathbf{r}} = m(r\hat{\mathbf{r}}) \times (\dot{r}\hat{\mathbf{r}} + r\dot{\theta}\hat{\theta})$$

Because $\hat{\mathbf{r}} \times \hat{\mathbf{r}} = 0$ and $|\hat{\mathbf{r}} \times \hat{\theta}| = 1$, we get the expressions

$$\mathbf{L} = mr^2\dot{\theta} (\hat{\mathbf{r}} \times \hat{\theta}) \implies |L| = mr^2\dot{\theta}$$

we see that $r^2\dot{\theta}$ represents magnitude of angular momentum per unit mass. We will denote this quantity l :

$$l = r^2\dot{\theta}$$

- Returning to the radial component of the equation of motion (corresponding to the $\hat{\mathbf{r}}$ terms) we have

$$m(\ddot{r} - r\dot{\theta}^2) = -\frac{dV}{dr}$$

Because $l = r^2\dot{\theta}$ is conserved, we can write $\dot{\theta}^2 = \frac{l^2}{r^4}$ and get

$$m\ddot{r} = -\frac{dV}{dr} + \frac{ml^2}{r^3}$$

But this is excellent! We have reduced the central force problem, originally in three dimensions, to a one-dimensional problem of radial motion. This is possible because the angular momentum vector is conserved: earlier, conservation of direction reduced our problem to planar motion, and now conservation of magnitude further reduced the problem to one dimension.

3.1.5 Effective Potential

- In the central force problem it is conventional (and convenient) to define an *effective potential* constructed to simplify the equation of motion. This is:

$$V_{\text{eff}} = V(r) + \frac{ml^2}{2r^2} \quad \implies \quad m\ddot{r} = -\frac{dV_{\text{eff}}}{dr}$$

The term $\frac{ml^2}{2r^2}$ is called the angular momentum or centrifugal barrier. It prevents a particle from getting too close to the origin, and arises because of orbital motion. We'll get a better feel for it in the next few paragraphs.

- In three dimensions, the system's total energy (the sum of the kinetic and potential energies) is

$$\begin{aligned} E &= \frac{1}{2}m\dot{\mathbf{r}}^2 + V(r) = \frac{1}{2}m\left(\dot{r}\hat{\mathbf{r}} + r\dot{\theta}\hat{\boldsymbol{\theta}}\right)^2 + V(r) \\ &= \frac{1}{2}m\dot{r}^2 + \frac{1}{2}mr^2\dot{\theta}^2 + V(r) \end{aligned}$$

where the second equality holds because $\hat{\mathbf{r}}$ and $\hat{\boldsymbol{\theta}}$ are orthonormal. Using the conservation of l which implies $\dot{\theta}^2 = \frac{l^2}{r^4}$ gives

$$E = \frac{1}{2}m\dot{r}^2 + \frac{ml^2}{2r^2} + V(r) = \frac{1}{2}m\dot{r}^2 + V_{\text{eff}}(r)$$

This result is important: we have derived the same effective potential as the one-dimensional problem, but began the analysis in three dimensions. This basically says the effective one dimensional system governed by $m\ddot{r} = -\frac{dV_{\text{eff}}}{dr}$ coincides with the three dimensional system.

We also see the extra term in the effective potential is a contribution from the angular kinetic energy.

3.1.6 Orbits with $V(r) = -\frac{k}{r}$ Classified by energy

We now turn to a concrete example: we will study the classic $\frac{1}{r}$ potential $V(r) = -\frac{k}{r}$. In this case the effective potential is

$$V_{\text{eff}} = -\frac{k}{r} + \frac{ml^2}{2r^2}$$

The minimum of the effective potential is:

$$r_{\min} = \frac{ml^2}{k} \quad \text{and} \quad V_{\min} = V(r_{\min}) = -\frac{k^2}{2ml^2}$$

We can use these results to classify the possible forms of motion in a $\frac{1}{r}$ potential based on the system's energy E (it helps to graph the effective potential as a function of r). There are three options

1. If $E = V_{\min}$ the particle rests at the minimum of the potential well V_{eff} and stays at the distance r_{\min} from the origin. Although the particle has fixed radial position, it still moves with angular speed $\dot{\theta} = \frac{l}{r_{\min}^2}$: the particle moves in a circular orbit about the origin.
2. If $E \in (V_{\min}, 0)$, the particle is trapped in a potential well and oscillates back and forth between the points of maximum energy. Since the particle also has angular speed $\dot{\theta} = \frac{l}{r_{\min}^2}$, the particle moves in a closed orbit in which radial distance r varies with time. For $V = -\frac{k}{r}$, this turns out to be an ellipse.
3. If $E > 0$, the particle has enough energy to overcome the potential well so the corresponding is unbounded. The particle approaches the origin from infinity, reaches a point of closest approach, and returns back out to infinity. For $V = -\frac{k}{r}$, the exact trajectory turns out to be a hyperbola.

3.1.7 Existence and Stability of Circular Orbits

We now consider a general central potential $V = V(r)$ and ask when can circular orbits exist, and when are they stable?

- Circular orbits exist when there is a solution of the equation $m\ddot{r} = -\frac{dV_{\text{eff}}}{dr}$ with $\dot{r} = 0$ (constant radius) and fixed $l \neq 0$ (non-zero angular speed) for all time. The condition $\dot{r} = 0$ obviously implies $\ddot{r} = 0$, so the equation of motion reads

$$0 \equiv m\ddot{r} = -\frac{dV_{\text{eff}}}{dr} \implies \frac{dV_{\text{eff}}}{dr} = 0$$

In other words, circular orbits correspond to critical points of the effective potential V_{eff} .

- The orbit is stable if small perturbations return the particle to the critical point, which means the critical point must be a minimum of V_{eff} (i.e. $V_{\text{eff}}'' > 0$) for a stable orbit.
- Let r_* denote the critical point. In terms of the original potential $V(r)$, the circular orbit condition $V_{\text{eff}}(r_*) = 0$ reads

$$V'(r_*) = \frac{ml^2}{r_*^3}$$

(recall $V_{\text{eff}} = V(r) + \frac{ml^2}{2r^2}$). Meanwhile stability condition $V''(r_*) > 0$ reads

$$V''(r_*) + \frac{3ml^2}{r_*^4} > 0$$

Finally, in terms of the central force $F(r) = -V'(r)$, the stability condition is

$$F'(r_*) + \frac{3}{r_*}F(r_*) < 0$$

3.2 The Orbit Equation

In this section we study the trajectories of particles under central forces, which we call *orbits*.

- It is straightforward to find how radial position $r(t)$ changes with time. Conservation of energy comes to the rescue. Since E is a constant of motion, we can write

$$E = \frac{1}{2}m\dot{r}^2 + V_{\text{eff}}(r)$$

and then invert the expression to give

$$\frac{dr}{dt} = \sqrt{\frac{2}{m}(E - V_{\text{eff}}(r))}$$

This could then be integrated, numerically if nothing else, to find an expression for $r(t)$.

When classifying orbits, it turns out to be more useful to find radial position as a function of the polar angle θ ; we then get a sense of the orbit by finding the shape of $r(\theta)$.

- We start our quest for $r(\theta)$ with two tricks. The first is to temporarily replace r with the new coordinate

$$u = \frac{1}{r}$$

Why? It just works: the math becomes much easier. I don't know of a good physical explanation. The second trick is our good friend the chain rule used with $\dot{\theta} = \frac{l}{r^2}$.

- We start with an expression for radial velocity \dot{r}

$$\frac{dr}{dt} = \frac{dr}{d\theta} \frac{d\theta}{dt} = \frac{dr}{d\theta} \frac{l}{r^2} = -l \frac{du}{d\theta}$$

where we have used $du = -\frac{dr}{r^2}$. Meanwhile, radial acceleration is

$$\frac{d^2r}{dt^2} = -l \frac{d}{dt} \left[\frac{du}{d\theta} \right] = -l \frac{d^2u}{d\theta^2} \dot{\theta} = -\frac{l^2}{r^2} \frac{d^2u}{d\theta^2} = -l^2 u^2 \frac{d^2u}{d\theta^2}$$

where the $\dot{\theta}$ term comes from the chain rule and $u = u(\theta)$.

- Recall the original radial orbit equation

$$m\ddot{r} = -\frac{dV}{dr} + \frac{ml^2}{r^3}$$

In terms of u , using our expression for \ddot{r} , the equation reads

$$\frac{d^2u}{d\theta^2} + u = \frac{1}{ml^2u^2} \left. \frac{dV}{dr} \right|_{r=\frac{1}{u}}$$

This is *orbit equation*, which we want to solve for $u(\theta)$, which would then give us $r(\theta)$. We could also extract the time dependence of r with $l = r^2\dot{\theta}$.

3.2.1 The Kepler Problem

- The *Kepler problem* is just the name given to understanding planetary orbits about a star under the influence of gravity. The correct potential for this scenario is

$$V_{\text{grav}}(r) = -\frac{GmM}{r}$$

where m and M are the mass of the planet and sun, respectively, while G is the universal gravitational constant. However, we will be a bit more general and work with the potential

$$V(r) = -\frac{km}{r}$$

with the tacit understanding that $k = GM$. Why? By using the general form, the results we derive will apply equally well to the electrostatic potential; we just have replace $k = \frac{-qQ}{4\pi\epsilon_0 m}$.

- For the potential $V = -\frac{km}{r}$, the orbit equation is simply

$$\frac{d^2u}{d\theta^2} + u = \frac{k}{l^2}$$

This is the equation for a harmonic oscillator with its center displaced by $\frac{k}{l^2}$. We typically write the solution to this equation in the form $u = A \cos \theta + B \sin \theta + \frac{k}{l^2}$, but for the orbit problem, it is much more useful to write

$$u(\theta) = A \cos(\theta - \theta_0) + \frac{k}{l^2}$$

where the A and θ_0 are integration constants.

- We're almost there. At the orbit's point of closest approach, r is smallest and u is largest. From our solution $u(\theta) = A \cos(\theta - \theta_0) + \frac{k}{l^2}$, we see that u is largest when the cosine term equals one, meaning the point of closest approach occurs when $\theta - \theta_0 = 0$.

This result provides a natural choice for our polar coordinate system: we orient our axes so the point of closest approach occurs at $\theta = 0$. To satisfy $\theta - \theta_0 = 0$ we then have the simple result $\theta_0 = 0$. We then have the solution

$$u(\theta) = A \cos \theta + \frac{k}{l^2} \quad \text{or} \quad r(\theta) = \frac{1}{A \cos \theta + \frac{k}{l^2}}$$

- By convention, we re-write the solution for $r(\theta)$ in the form

$$r(\theta) = \frac{r_0}{\epsilon \cos \theta + 1}$$

where $r_0 = \frac{l^2}{k}$ and $\epsilon = \frac{Al^2}{k}$ (we simply multiplied $r(\theta)$ above and below by $\frac{l^2}{k}$).

Note that r_0 is not an integration constant; it is uniquely determined by the fixed angular momentum l . The *eccentricity* ϵ is effectively the new integration constant.

3.2.2 Analyzing the Solutions

Note that the solutions for $r(\theta)$ given by

$$r(\theta) = \frac{r_0}{\epsilon \cos \theta + 1}$$

are *conic sections* (e.g. ellipses, hyperbola, etc...). We can classify the orbits $r(\theta)$ by the eccentricity ϵ . There are a few choices:

1. For $\epsilon < 1$, the radial position is bounded in the interval

$$\frac{r_0}{r} \in [1 - \epsilon, 1 + \epsilon]$$

since $\cos \theta$ falls between -1 and 1 . This type of orbit is an ellipse. To double check, we can convert to Cartesian coordinates with the substitutions $x = r \cos \theta$, $y = r \sin \theta$ and $r^2 = x^2 + y^2$. We have

$$r = \frac{r_0}{\epsilon \cos \theta + 1} \implies r^2 = (r_0 - \epsilon r \cos \theta)^2 \implies x^2 + y^2 = (r_0 - \epsilon x)^2$$

A few more steps of algebra lead to the Cartesian equation of an ellipse

$$\frac{(x - x_0)^2}{a^2} + \frac{y^2}{b^2} = 1$$

where

$$x_0 = -\frac{\epsilon r_0}{1 - \epsilon^2} \quad a^2 = \frac{r_0^2}{(1 - \epsilon^2)^2} \quad b^2 = \frac{r_0^2}{1 - \epsilon^2}$$

The two axes of the ellipse have lengths a and b with $a > b$. The center of gravitational attraction (e.g. the sun) sits at the origin $r = 0$; the origin and ellipse's center are separated by

$$|x_0| = \frac{\epsilon r_0}{1 - \epsilon^2} = \epsilon a$$

The origin of gravitational attraction is a *focus* of the ellipse; the other focus lies on the major axis an equal distance from the ellipse's center. Note that when $\epsilon = 0$, the focus and center coincide and the two axes have equal lengths $a = b$: this is a circular orbit.

2. Orbits with $\epsilon > 1$ are hyperbolae. Repeating analogous algebraic steps to the elliptical orbit leads to the Cartesian coordinate equation

$$\frac{1}{a^2}(x + x_0)^2 - \frac{y^2}{b^2} = 1$$

with a, b and x_0 defined as before.

The hyperbola's asymptotes make angles $\cos \theta = -\frac{1}{\epsilon}$ with the x axis. This should make sense; from the polar orbit equation $r(\theta)$ we see that $r \rightarrow \infty$ when the denominator $\epsilon \cos \theta + 1$ equals zero; this occurs for $\cos \theta = -\frac{1}{\epsilon}$.

3. The special case $\epsilon = 1$ is particularly simple: it is a hyperbolic orbit. The Cartesian coordinate equation is

$$y^2 = r_0^2 - 2r_0x$$

3.2.3 Energy of the Orbit Revisited

- The solution of the orbit equation $r(\theta) = \frac{r_0}{\epsilon \cos \theta + 1}$ gives us an interesting new look at the orbit energy. Recall that

$$E = \frac{1}{2}m\dot{r}^2 + V_{\text{eff}}(r) = \frac{1}{2}m\dot{r}^2 - \frac{km}{r} + \frac{ml^2}{2r^2}$$

re-writing \dot{r} with the chain rule and substituting $\dot{\theta} = \frac{l}{r^2}$ gives

$$\begin{aligned} E &= \frac{1}{2}m \left(\frac{dr}{d\theta} \right)^2 \dot{\theta}^2 - \frac{km}{r} + \frac{ml^2}{2r^2} \\ &= \frac{1}{2}m \left(\frac{dr}{d\theta} \right)^2 \frac{l^2}{r^4} - \frac{km}{r} + \frac{ml^2}{2r^2} \end{aligned}$$

- We can differentiate our solution $r(\theta) = \frac{r_0}{\epsilon \cos \theta + 1}$ to get

$$\frac{dr}{d\theta} = \frac{r_0 \epsilon \sin \theta}{(1 + \epsilon \cos \theta)^2}$$

Plugging this in to the energy expression and a few steps of rather satisfying algebra (using $\cos^2 \theta + \sin^2 \theta = 1$ and $r_0 = \frac{l^2}{k}$) leads to the simple result

$$E = \frac{mk^2}{2l^2}(\epsilon^2 - 1)$$

Notice that all θ and r dependence has vanished from the expression for E . This makes sense: energy is a constant of motion.

- With our nice result $E = \frac{mk^2}{2l^2}(\epsilon^2 - 1)$ we now return the three cases of orbit classified by eccentricity to see that:
 1. Orbits with $\epsilon < 1$ have energy $E < 0$. Negative energy means a bounded orbit, in agreement with our earlier conclusion that orbits $\epsilon < 1$ are ellipses.
 2. Orbits with $\epsilon > 1$ have positive energy $E > 0$; these are unbounded orbits, in agreement with our earlier result that $\epsilon > 1$ leads to hyperbolic orbits.
 3. Finally, $\epsilon = 0$ leads to $E = -\frac{mk^2}{2l^2}$. We have seen this before: it is the minimum of the effective potential, and we know from earlier that orbits with minimum V_{eff} are circular, in agreement with our geometric interpretation that ellipses with $\epsilon = 0$ are circles.

3.2.4 Kepler's Laws of Planetary Motion

Kepler's three laws are obeyed by all planets in the solar system and were first published by Kepler in 1605.

- Kepler's first law states:

Each planet moves in an ellipse, with the Sun at one focus.

Without knowing, we have dedicated a good part of the past few sections to Kepler's first law when discussing the elliptical solution to the orbit equation (when $\epsilon < 1$).

- Kepler's second law is:

The line between the a planet and the Sun sweeps out equal areas in equal times at any point in the orbit.

The second law follows directly from conservation of angular momentum. From the equation of area in polar coordinates ($A = \frac{1}{2} \int r(\theta)^2 d\theta$), we that in the time δt an orbiting planet sweeps out the area

$$\delta A = \frac{1}{2} r^2 \delta \theta$$

Differentiating both sides with respect to time leads to

$$\frac{dA}{dt} = \frac{1}{2} r^2 \dot{\theta} = \frac{l}{2}$$

But we know that l is constant by conservation of angular momentum, so $\frac{dA}{dt}$ is also constant throughout the orbit. Because conservation of l holds for all central force motion, Kepler's second law would hold for any central potential $V = V(r)$, not just the gravitational potential $V = -\frac{GMm}{r}$.

- Kepler's third law states:

The period of orbit is proportional the $\frac{3}{2}$ power of the radius, i.e. $T \sim r^{\frac{3}{2}}$.

We turn to the inverse square law of gravitation

$$F = \frac{GMm}{r^2}$$

There is a neat proof of the third law that follows simply from dimensional analysis. Kepler's third law is essentially a relation ship between a length and time describing the orbit (e.g. an average radius R and the period T). We recognize that the only parameter in the problem is GM , so the radius and period in Kepler's law must be related by GM .

Because GM has dimensions $\text{length}^3 \cdot \text{time}^{-2}$, or

$$[GM] = L^3 T^{-2}$$

By dimensional analysis, Kepler's law relating R, T and GM must then be of the form

$$GM \sim \frac{R^3}{T^2} \implies T^2 \sim \frac{R^3}{GM} \quad \text{or} \quad T \sim R^{\frac{3}{2}}$$

- For an elliptical orbit we can be more precise. Using the area of an ellipse $A = \pi ab$ and our earlier relationships between a, b and r_0 we have

$$A = \pi ab = \pi a^2 \sqrt{1 - \epsilon^2} = \frac{\pi r_0^2}{(1 - \epsilon^2)^{3/2}}$$

Next, we use Kepler's second law: since area is swept out at a constant rate $\frac{dA}{dt} = \frac{l}{2}$, we have $\frac{A}{T} = \frac{l}{2}$ or

$$T = \frac{2A}{l} = \frac{2\pi r_0^2}{l(1 - \epsilon^2)^{3/2}} = \frac{2\pi}{\sqrt{GM}} \left(\frac{r_0}{1 - \epsilon^2} \right)^{3/2}$$

We see that T is indeed proportional to a radius quantity squared. But what exactly does the quantity $\frac{r_0}{1-\epsilon^2}$ represent? It turns out to have a very nice interpretation. From the orbit equation $r(\theta)$ we see that the points of minimum and maximum approach occur at

$$r_{\min} = \frac{r_0}{(1 + \epsilon)} \quad \text{and} \quad r_{\max} = \frac{r_0}{(1 - \epsilon)}$$

respectively. The average of these quantities is none other than

$$R = \frac{1}{2}(r_{\min} + r_{\max}) = \frac{r_0}{(1 - \epsilon^2)}$$

R is the average radius of the orbit. In this interpretation, we have

$$T = \frac{2\pi}{\sqrt{GM}} R^{3/2}$$

and Kepler's third law emerges crystal clear.

3.2.5 The Two-Body Central Force Problem with Lagrangian Mechanics

A new approach to the same problem: solving for the motion of two bodies under the influence of a central force two-body problem. We'll see the Lagrangian approach gives the same result as the Newtonian one, as it must.

- As usual, we have our two bodies with position vectors and masses \mathbf{r}_1, m_1 and \mathbf{r}_2, m_2 , respectively. The center of mass \mathbf{R} is and reduced mass μ are

$$(m_1 + m_2)\mathbf{R} = m_1\mathbf{r}_1 + m_2\mathbf{r}_2 \quad \text{and} \quad \mu = \frac{m_1 m_2}{m_1 + m_2}$$

while the relative position vector \mathbf{r} and its magnitude r are

$$\mathbf{r} = \mathbf{r}_1 - \mathbf{r}_2 \quad \text{and} \quad r = |\mathbf{r}|$$

The system's Lagrangian is

$$L = T - V = \frac{1}{2} (m_1 \dot{\mathbf{r}}_1^2 + m_2 \dot{\mathbf{r}}_2^2) - V(r)$$

- We start by writing the Lagrangian in terms of the center of mass \mathbf{R} and relative position \mathbf{r} :

$$L = \frac{1}{2} (m_1 + m_2) \dot{\mathbf{R}}^2 + \frac{1}{2} \mu \dot{\mathbf{r}}^2 - V(r)$$

Note that the Lagrangian decomposes into a piece describing the center of mass \mathbf{R} and a piece describing the relative separation \mathbf{r} .

Since there are no external forces acting on the system the motion of the center of mass is uniform, i.e. $\dot{\mathbf{R}}$. This means the $\frac{1}{2}(m_1 + m_2)\dot{\mathbf{R}}^2$ term is also constant and thus won't affect the system's dynamics.

- We know from the Newtonian approach that the angular momentum \mathbf{L} is conserved; it follows trivially from $\dot{\mathbf{L}} = \mathbf{r} \times \mathbf{F} = 0$ for central forces. Here's an elegant Lagrange-style approach to the same result: our system's Lagrangian is invariant under infinitesimal

rotations, so Noether's theorem tells us the system's angular momentum must be conserved. The angular momentum is

$$\mathbf{L} = \mathbf{r} \times \mathbf{p}$$

where $\mathbf{p} = \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{\mathbf{r}}} \right)$ is the momentum conjugate to \mathbf{r} ; because we're using Cartesian coordinates, \mathbf{p} coincides with the Newtonian momentum $m\dot{\mathbf{r}}$.

- Since $\mathbf{L} = \mathbf{r} \times \mathbf{p}$ is perpendicular to \mathbf{r} by definition, the system's orbit must lie in a plane perpendicular to \mathbf{L} . Planar motion is best suited to the polar coordinates (r, ϕ) , in terms of which the Lagrangian is

$$L = \frac{1}{2}\mu \left(\dot{r}^2 + r^2 \dot{\phi}^2 \right) - V(r)$$

We immediately see L is independent of ϕ , which means (see Section 2.4.1) that the quantity

$$\frac{\partial L}{\partial \dot{\phi}} = \mu r^2 \dot{\phi} \equiv J$$

is conserved. J is the magnitude of angular momentum $|\mathbf{L}|$; I'm denoting it by J and not L to avoid confusion with the Lagrangian.

- We get the system's equations of motion from the Lagrange equations for the coordinate r :

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{r}} \right) - \frac{\partial L}{\partial r} = \mu \ddot{r} - \mu r \dot{\phi}^2 + \frac{\partial V}{\partial r} = 0$$

We can then eliminate $\dot{\phi}$ by writing it in terms of the conserved quantity J

$$\dot{\phi}^2 = \frac{J^2}{\mu^2 r^4} \implies \mu \ddot{r} = \frac{J^2}{2\mu r^3} - V(r) \equiv -\frac{\partial}{\partial r} V_{\text{eff}}(r)$$

where V_{eff} is the effective potential

$$V_{\text{eff}}(r) = V(r) + \frac{J^2}{2\mu r^2}$$

- Next, since the Lagrangian is not an explicit function of time (i.e. $\frac{\partial L}{\partial t} = 0$), the system's energy is conserved (see 2.4.1) so

$$E = \frac{1}{2}\mu \dot{r}^2 + V_{\text{eff}}(r)$$

is a constant of motion. In principle, this is enough to solve the dynamics as long as V_{eff} is known. We can rearrange the above equation, separate variables, and integrate to get

$$\int dt = \sqrt{\frac{\mu}{2}} \int \frac{dr}{\sqrt{E - V_{\text{eff}}(r)}}$$

from which we could solve for $r(t)$, numerically if nothing else.

At this point we've reduced the problem to a one-dimensional problem with a single degree of freedom r , just like in the Newtonian approach, summed up in our main result

$$E = \frac{1}{2}\mu \dot{r}^2 + V_{\text{eff}}(r)$$

The procedure from here on is analogous to what we've already solved, so I won't repeat.

3.3 Laplace-Runge-Lenz Vector

3.3.1 Derivation

- Let's return to the two body central force problem, denoting the positions of the two bodies by \mathbf{r}_1 and \mathbf{r}_2 and defining the relative position vector \mathbf{r} as

$$\mathbf{r} = \mathbf{r}_1 - \mathbf{r}_2$$

Because the force is assumed to be central, Newton's second law can be written

$$m\ddot{\mathbf{r}} = F(r)\hat{\mathbf{r}}$$

where $\hat{\mathbf{r}}$ is the unit vector in the direction of the relative position vector and m is the reduced mass. For the gravitational force $F(r) = \frac{-k}{r^2}$ we have

$$m\ddot{\mathbf{r}} = \frac{-k}{r^2}\hat{\mathbf{r}}$$

- Next, we take right-side cross product of both sides of Newton's law with the angular momentum $\mathbf{L} = m\mathbf{r} \times \dot{\mathbf{r}}$. The result is

$$\ddot{\mathbf{r}} \times \mathbf{L} = k\dot{\hat{\mathbf{r}}}$$

To start we first compute the time derivative $\dot{\hat{\mathbf{r}}}$; this seems odd, but it turns out super useful in the next step. The result is

$$\dot{\hat{\mathbf{r}}} = \frac{\dot{\mathbf{r}}(\mathbf{r} \cdot \mathbf{r}) - (\mathbf{r} \cdot \dot{\mathbf{r}})\mathbf{r}}{r^3} = \frac{r^2\dot{\mathbf{r}} - (\mathbf{r} \cdot \dot{\mathbf{r}})\mathbf{r}}{r^3}$$

We then return to cross product of Newton's law to get

$$m\ddot{\mathbf{r}} \times \mathbf{L} = \frac{-k}{r^2}\hat{\mathbf{r}} \times \mathbf{L} \implies m\ddot{\mathbf{r}} \times \mathbf{L} = -\frac{km}{r^3}\mathbf{r} \times (\mathbf{r} \times \dot{\mathbf{r}})$$

We cancel masses, use the identity $\mathbf{a} \times (\mathbf{b} \times \mathbf{c}) = (\mathbf{a} \cdot \mathbf{c})\mathbf{b} - (\mathbf{a} \cdot \mathbf{b})\mathbf{c}$, and conveniently recognize the expression for $\dot{\hat{\mathbf{r}}}$:

$$\ddot{\mathbf{r}} \times \mathbf{L} = -\frac{k}{r^3}((\mathbf{r} \cdot \dot{\mathbf{r}})\mathbf{r} - r^2\dot{\mathbf{r}}) = +k\frac{(r^2\dot{\mathbf{r}} - (\mathbf{r} \cdot \dot{\mathbf{r}})\mathbf{r})}{r^3} = k\dot{\hat{\mathbf{r}}}$$

- Because $\dot{\mathbf{L}} = 0$ (i.e. \mathbf{L} is conserved in any central force problem), we can write the left hand side in the form $\ddot{\mathbf{r}} \times \mathbf{L} = \frac{d}{dt}[\dot{\mathbf{r}} \times \mathbf{L}]$ which then gives us

$$\frac{d}{dt}[\dot{\mathbf{r}} \times \mathbf{L}] = k\dot{\hat{\mathbf{r}}}$$

Since k is a constant, we can integrate the equation with respect to time to get

$$\dot{\mathbf{r}} \times \mathbf{L} = k\hat{\mathbf{r}} + \mathbf{x}$$

where $\mathbf{x} \neq \mathbf{x}(t)$ is a time-independent vector resulting from the integration over time. (You could think of \mathbf{x} as a vector integration constant).

- Finally, we divide the equation by k and rearrange to get

$$\mathbf{A} \equiv \frac{\mathbf{x}}{k} = \frac{\dot{\mathbf{r}} \times \mathbf{L}}{k} - \hat{\mathbf{r}}$$

where we have defined the Laplace-Runge-Lenz vector as $\mathbf{A} \equiv \frac{\mathbf{x}}{k}$. Because both \mathbf{x} and k are time independent so is \mathbf{A} , i.e. the LRL vector is conserved for the $F \sim \frac{1}{r^2}$ central force problem.

3.3.2 The LRL Vector and the Orbit Equation

The LRL vector gives us a neat way of deriving the orbit equation for the Kepler problem of planetary motion.

- Starting with the definition of the LRL vector, we take the dot product of both sides with \mathbf{r} . With the help of the vector identity $(\mathbf{a} \times \mathbf{b}) \cdot \mathbf{c} = (\mathbf{c} \times \mathbf{a}) \cdot \mathbf{b}$ we get

$$\mathbf{A} \cdot \mathbf{r} = \frac{\dot{\mathbf{r}} \times \mathbf{L}}{k} \cdot \mathbf{r} - \hat{\mathbf{r}} \cdot \mathbf{r} = \frac{(\mathbf{r} \times \dot{\mathbf{r}}) \cdot \mathbf{L}}{k} - r = \frac{\mathbf{L} \cdot \mathbf{L}}{mk} - r = \frac{L^2}{km} - r$$

- If we write $\mathbf{A} \cdot \mathbf{r} = Ar \cos \theta$ where θ is the angle between \mathbf{A} and \mathbf{r} , and solve for r we get

$$r = \frac{L^2}{km} \frac{1}{1 + A \cos \theta}$$

This should look familiar: it is our solution to the orbit equation for the Kepler problem with the perihelion set to $\theta = 0$. We had previously written the solution as

$$r(\theta) = \frac{r_0}{1 + \epsilon \cos \theta} \quad \text{where} \quad r_0 = \frac{l^2}{\tilde{k}}, \quad l = \frac{L}{m}, \quad \tilde{k} = Gm$$

- If you look closely, the k 's don't seem to match. Not to fear, it's just because we weren't using the reduced mass in our original solution of the orbit equation. Here's a detailed explanation; feel free to skip. I have explicitly denoted the old constant from our first encounter with the orbit equation as $\tilde{k} = Gm$. This time around (recall we wrote $F = -\frac{k}{r^2}$) we have been using $k = GMm$ where M is the mass of the large body and m the mass of the smaller, orbiting body. We're left in a notational muddle, since we have also used m for the reduced mass $\frac{mM}{m+M}$. But in the typical Kepler problem with $M \gg m$ the reduced mass is nearly equal to the mass of the smaller body. We then have $k = m\tilde{k}$, and we see both orbit equations agree.
- Comparing the two orbit equations reveals an awesome result: comparing the coefficient of the $\cos \theta$ term tells us

$$A = \epsilon$$

In other words, the eccentricity of the orbit is the magnitude of the LRL vector.

3.4 Scattering

TODO

4 The Motion of Rigid Bodies

- Rigid bodies are objects (discrete or continuous distributions of mass points) where the distance between the points is fixed, i.e. $|\mathbf{r}_i - \mathbf{r}_j|$ is constant for all i, j . Rigid bodies have zero internal degrees of freedom because their constituent points are fixed.
- In general, a rigid body with a free center of mass has six degrees of freedom: three from rotation and three from translation. A rigid body with a *fixed* center of mass loses its translational freedom, so it has only three rotational degrees of freedom.

The general motion of a free rigid body can be separated into a translation of the center of mass and rotation about a point, often chosen to be the center of mass.

4.1 Rigid Body Kinematics

4.1.1 On Frames and Degrees of Freedom

- In general the analysis involves a fixed (inertial) lab frame $\{\tilde{\mathbf{e}}_a\}$ and a body-fixed frame $\{\mathbf{e}_a\} = \{\mathbf{e}_a(t)\}$ that moves with the body. Both frames have orthonormal axes: $\tilde{\mathbf{e}}_a \cdot \tilde{\mathbf{e}}_b = \delta_{ab}$ and $\mathbf{e}_a(t) \cdot \mathbf{e}_b(t) = \delta_{ab}$.
- For all times t , there exists a *unique* rotation matrix $\mathbf{R}(t)$ for which $\mathbf{e}_a(t) = \mathbf{R}_{ab}(t)\tilde{\mathbf{e}}_b$. In other words, it is always possible to rotate the lab frame in such a way that the directions of its axes $\{\tilde{\mathbf{e}}_a\}$ coincide with the body-fixed frame $\{\mathbf{e}_a\}$, and the corresponding rotation matrix \mathbf{R} is unique.
- As the rigid body rotates, we can thus describe its orientation with the 3×3 time-dependent rotation matrix $\mathbf{R}(t)$. As for any rotation matrix \mathbf{R} is orthogonal, $\mathbf{R}^T \mathbf{R} = \mathbf{I}$ and $\det \mathbf{R} = 1$. In general, a 3×3 matrix has 9 degrees of freedom, but the orthogonality condition $\mathbf{R}^T \mathbf{R} = \mathbf{I}$ imposes 6 constraints, leading to the aforementioned 3 degrees of rotational freedom. We thus need three coordinates to parameterize a rigid body's rotation.

4.1.2 Angular Velocity

I hope you're comfortable with indexes and Einstein summation convention. There's a healthy dose of it coming fast.

- Any point in the body $\mathbf{r}(t)$ can be expanded in either the body frame or space frame:

$$\mathbf{r}(t) = r_a \mathbf{e}_a(t) = \tilde{r}_a(t) \tilde{\mathbf{e}}_a$$

where $\tilde{r}_a(t) = r_a R_{ab}(t)$. In the body-fixed frame, the distances r_a are constant and in the lab frame, the basis vectors $\tilde{\mathbf{e}}_a$ are constant.

- Taking the time derivative of $\mathbf{r}(t)$ produces

$$\begin{aligned} \dot{\mathbf{r}}(t) &= \tilde{\mathbf{e}}_a \frac{d\tilde{r}_a}{dt} && \text{(lab frame)} \\ &= r_a \frac{d\mathbf{e}_a}{dt} && \text{(body-fixed frame)} \\ &= r_a \frac{dR_{ab}}{dt} \tilde{\mathbf{e}}_b && \text{(using } \tilde{r}_b = r_a R_{ab}) \end{aligned}$$

- Meanwhile, using the transformation $\mathbf{e}_b = \mathbf{R}_{ab} \tilde{\mathbf{e}}_a$ and the fact that the lab basis vectors $\tilde{\mathbf{e}}_a$ are constant, we can calculate the derivative of the body-fixed basis vectors $\{\mathbf{e}_a(t)\}$ to be

$$\frac{d\mathbf{e}_a}{dt} = \frac{d}{dt}[\mathbf{R}_{ab} \tilde{\mathbf{e}}_b] = \frac{d\mathbf{R}_{ab}}{dt} \tilde{\mathbf{e}}_b$$

Next, using the transformation $\tilde{\mathbf{e}}_b = \mathbf{R}_{bc}^{-1} \mathbf{e}_c$, and \mathbf{R} 's orthogonality $\mathbf{R}_{ij}^{-1} = \mathbf{R}_{ji}$ we finish the calculation with

$$\frac{d\mathbf{e}_a}{dt} = \dots = \frac{d\mathbf{R}_{ab}}{dt} \tilde{\mathbf{e}}_b = \frac{d\mathbf{R}_{ab}}{dt} \mathbf{e}_c \equiv \Omega_{ac} \mathbf{e}_c$$

where we've defined $\Omega_{ac} \equiv \frac{d\mathbf{R}_{ab}}{dt} \mathbf{R}_{cb}$, the elements of a 3×3 matrix $\mathbf{\Omega}$.

The matrix $\mathbf{\Omega}$ defined above is antisymmetric, i.e. $\Omega_{ac} = -\Omega_{ca}$. To show this, we differentiate the identity $\mathbf{R}_{ab} \mathbf{R}_{cb} = \delta_{ac}$ and recognize the definition of Ω_{ac}

$$\mathbf{R}_{ab} \mathbf{R}_{cb} = \delta_{ac} \implies \frac{d\mathbf{R}_{ab}}{dt} \mathbf{R}_{cb} + \mathbf{R}_{ab} \frac{d\mathbf{R}_{cb}}{dt} = 0 \implies \Omega_{ac} + \Omega_{ca} = 0$$

Because $\mathbf{\Omega}$ is antisymmetric (its diagonal terms are necessarily zero and its off-diagonal terms are related by $\Omega_{ab} = -\Omega_{ba}$) it has only 3 independent components. To be concrete, $\mathbf{\Omega}$ looks something like this:

$$\mathbf{\Omega} = \begin{bmatrix} 0 & \Omega_{12} & \Omega_{13} \\ \Omega_{21} & 0 & \Omega_{23} \\ \Omega_{31} & \Omega_{32} & 0 \end{bmatrix} = \begin{bmatrix} 0 & \Omega_{12} & -\Omega_{31} \\ -\Omega_{12} & 0 & \Omega_{23} \\ \Omega_{31} & -\Omega_{23} & 0 \end{bmatrix}$$

Since $\mathbf{\Omega}$ has only 3 independent components, we can define an equivalent three-component, one-index object (i.e. a vector) $\boldsymbol{\omega}$ with the Levi-Civita symbol with the formula

$$\omega_a = \frac{1}{2} \epsilon_{abc} \Omega_{bc} \quad a, b, c = 1, 2, 3$$

By components, we get $\omega_1 = \Omega_{23}$, $\omega_2 = \Omega_{31}$ and $\omega_3 = \Omega_{12}$ (try it yourself!).

The ω_a form the components of a vector in the body-fixed frame $\{\mathbf{e}_a\}$: $\boldsymbol{\omega} = \omega_a \mathbf{e}_a$. In terms of $\boldsymbol{\omega}$, the time derivative of the body frame is

$$\dot{\mathbf{e}}_a(t) = \Omega_{ac} \mathbf{e}_c = -\epsilon_{abc} \omega_b \mathbf{e}_c = \boldsymbol{\omega} \times \mathbf{e}_a$$

where $\mathbf{e}_a \times \mathbf{e}_b = \epsilon_{abc} \mathbf{e}_c$. After all that work, we finally have an expression for $\dot{\mathbf{e}}(t)$. But it was worth it: along the way we have defined something very important, which we haven't yet discussed: the vector $\boldsymbol{\omega}$ is the body's *instantaneous angular velocity*. I'll stress again that its components ω_a are measured with respect to the body-fixed frame $\{\mathbf{e}_a\}$. Angular velocity comes up everywhere in rigid body mechanics, we'll see much more of it in the next sections.

- Informal interpretation: the angular velocity encodes the speed and direction of the body's rotation. To see this, consider a point \mathbf{r} in the body, and rotate the point by the infinitesimal angle $d\phi$ about the fixed axis $\hat{\mathbf{n}}$, where the angle between \mathbf{r} and $\hat{\mathbf{n}}$ is θ .

The vector \mathbf{r} changes in magnitude by $|d\mathbf{r}| = |\mathbf{r}| \sin \theta d\phi$. Because the body is rigid, $d\mathbf{r} \perp \mathbf{r}$, so $d\mathbf{r} = d\phi \times \mathbf{r}$, where $d\phi = d\phi \hat{\mathbf{n}}$. "Dividing" $d\mathbf{r} = d\phi \times \mathbf{r}$ by dt produces

$$\frac{d\mathbf{r}}{dt} = \frac{d\phi}{dt} \times \mathbf{r} = \boldsymbol{\omega} \times \mathbf{r}$$

Thus, $\boldsymbol{\omega}$ is the body's instantaneous angular velocity about the rotation axis $\hat{\mathbf{n}}$. In general, both $\hat{\mathbf{n}}$ and $\boldsymbol{\omega}$ change over time.

- As an aside, much of rigid body dynamics involves determining the angular velocity vector $\boldsymbol{\omega}(t)$ of various rigid bodies as they spin and turn in space. Let's say we find $\boldsymbol{\omega}$. How do we get back to our starting point, the rotation matrix $\mathbf{R}(t)$, which describes the rotating body's orientation in space?

First, we expand the angular velocity in the body-fixed frame to get

$$\boldsymbol{\omega} = \omega_a \mathbf{e}_a$$

We then turn the vector into the 3×3 antisymmetric matrix $\boldsymbol{\Omega}$ via

$$\Omega_{ab} = \epsilon_{abc} \omega_c$$

Finally, we recover the rotation matrix \mathbf{R} by solving the differential equation

$$\boldsymbol{\Omega} = \frac{d\mathbf{R}}{dt} \mathbf{R}^{-1}$$

which comes from the definition of $\boldsymbol{\Omega}$ via $\Omega_{ac} = \frac{dR_{ab}}{dt} R_{cb} = \frac{dR_{ab}}{dt} R_{bc}^{-1}$. Note that this is a *matrix* differential equation. It is solved with a mathematical tool called path-ordered exponentials, which is beyond the scope of this course.

4.1.3 Inertia Tensor

- The inertia tensor arises naturally from computing the kinetic energy of a fixed rotating body:

$$\begin{aligned} T &= \frac{1}{2} \sum_i m_i \dot{\mathbf{r}}_i \cdot \dot{\mathbf{r}}_i = \frac{1}{2} \sum_i m_i (\boldsymbol{\omega} \times \mathbf{r}_i) \cdot (\boldsymbol{\omega} \times \mathbf{r}_i) \\ &= \frac{1}{2} \sum_i m_i [(\boldsymbol{\omega} \cdot \boldsymbol{\omega})(\mathbf{r}_i \cdot \mathbf{r}_i) - (\boldsymbol{\omega} \cdot \mathbf{r}_i)^2] \end{aligned}$$

where the last equality uses Lagrange's vector identity $(\mathbf{a} \times \mathbf{b})^2 = \mathbf{a}^2 \mathbf{b}^2 - (\mathbf{a} \cdot \mathbf{b})^2$. The equality can be written

$$T = \frac{1}{2} \omega_a I_{ab} \omega_b = \frac{1}{2} \boldsymbol{\omega} \cdot \mathbf{I} \boldsymbol{\omega}$$

where \mathbf{I} is the *inertia tensor*. Its components, *measured in the body-fixed frame*, are

$$I_{ab} = \sum_i m_i ((\mathbf{r}_i \cdot \mathbf{r}_i) \delta_{ab} - r_{ia} r_{ib})$$

The inertia tensor is always measured with respect to a particular point in space, which determines the origin for the position vectors \mathbf{r}_i . In general, the inertia tensor is different when measured with respect to different points in space.

Two important notes: By construction $I_{ab} = I_{ba}$, so \mathbf{I} is symmetric. Second, since \mathbf{I} is measured in the body-fixed frame, its components are independent of time.

For a continuous mass distribution with density $\rho = \rho(\mathbf{r})$, the inertia tensor is

$$\mathbf{I} = \int d^3 \mathbf{r} \rho(\mathbf{r}) \begin{bmatrix} y^2 + z^2 & -xy & -xz \\ -xy & x^2 + z^2 & -yz \\ -xz & -yz & x^2 + y^2 \end{bmatrix}$$

In principle, \mathbf{I} can be calculated with respect to any point in space.

- Because \mathbf{I} is real and symmetric, it can be diagonalized, i.e. there exists an orthogonal matrix \mathbf{Q} such that $\mathbf{I}' = \mathbf{Q}\mathbf{I}\mathbf{Q}^T$ is diagonal.

Equivalently, we can rotate the orientation of the body-fixed frame $\{\mathbf{e}_a\}$ to coincide with \mathbf{I} 's eigenvectors, which are given by $\{\mathbf{Q}\mathbf{e}_a\}$. In the frame $\{\mathbf{Q}\mathbf{e}_a\}$ the inertia tensor \mathbf{I} is diagonal.

The frame in which \mathbf{I} is diagonal is called the *frame of principle axes*. In this basis corresponding to the principle axes frame, \mathbf{I} takes the form

$$\mathbf{I} = \begin{bmatrix} I_1 & 0 & 0 \\ 0 & I_2 & 0 \\ 0 & 0 & I_3 \end{bmatrix}$$

\mathbf{I} 's eigenvalues $I_{1,2,3}$ are called the *principle moments of inertia*. By the way, the principle moments of inertia are always real and positive.

4.1.4 Parallel Axis Theorem

The parallel axis theorem is a handy tool for calculating inertia tensors. If we know a body's inertia tensor *about its center of mass*, the parallel axis theorem makes it relatively easy calculate the inertia tensor about any other point. Here's a mathematical formulation:

If $\mathbf{I}^{(\text{cm})}$ is a rigid body's inertia tensor measured with respect to the body's center of mass $\mathbf{r}^{(\text{cm})}$, then the inertia tensor \mathbf{I}' measured about some other point \mathbf{r}' displaced from $\mathbf{r}^{(\text{cm})}$ by the vector \mathbf{c} is given by

$$I'_{ab} = I_{ab}^{(\text{cm})} + M((\mathbf{c} \cdot \mathbf{c})\delta_{ab} - c_a c_b)$$

where M is the body's mass.

Proof: Starting from the definition of I'_{ab} and multiplying out:

$$\begin{aligned} I'_{ab} &= \sum_i m_i [((\mathbf{r}_i - \mathbf{c}) \cdot (\mathbf{r}_i - \mathbf{c}))\delta_{ab} - (r_i - c)_a (r_i - c)_b] \\ I'_{ab} &= \sum_i m_i [(\mathbf{r}_i \cdot \mathbf{r}_i)\delta_{ab} - r_{i_a} r_{i_b} + \{-2\mathbf{r}_i \cdot \mathbf{c}\delta_{ab} + r_{i_a} c_b + r_{i_b} c_a\} + ((\mathbf{c} \cdot \mathbf{c})\delta_{ab} - c_a c_b)] \\ &= I_{ab}^{(\text{cm})} + 0 + \sum_i m_i (\mathbf{c}^2 \delta_{ab} - c_a c_b) = I_{ab}^{(\text{cm})} + M(\mathbf{c}^2 \delta_{ab} - c_a c_b) \end{aligned}$$

The terms in the curly brackets $\{\dots\}$ in the middle line are linear in \mathbf{r}_i , so they vanish by definition of center of mass $\sum_i m_i \mathbf{r}_i = 0$. The additional term $M(\mathbf{c}^2 \delta_{ab} - c_a c_b)$ has an interesting physical interpretation: it is the inertia tensor (still measured with respect to \mathbf{r}') we'd find if the body's entire mass were concentrated into a point at the center of mass.

4.1.5 Angular Momentum

- Angular momentum, like angular velocity and the inertia tensor, is calculated with respect to a given point. In terms of the inertia tensor, using the vector triple product, we have:

$$\begin{aligned} \mathbf{L} &= \sum_i m_i \mathbf{r}_i \times \dot{\mathbf{r}}_i = \sum_i m_i \mathbf{r}_i \times (\boldsymbol{\omega} \times \mathbf{r}_i) \\ &= \sum_i m_i (r_i^2 \boldsymbol{\omega} - (\mathbf{r}_i \cdot \boldsymbol{\omega}) \mathbf{r}_i) \\ &= \mathbf{I} \boldsymbol{\omega} \end{aligned}$$

There is an important concept hiding in the result $\mathbf{L} = \mathbf{I}\boldsymbol{\omega}$: in general angular velocity $\boldsymbol{\omega}$ and angular momentum \mathbf{L} point in different directions.

There are probably more elegant ways of showing this, but this way worked for me. By components, $\mathbf{L} = \mathbf{I}\boldsymbol{\omega}$ reads

$$\begin{bmatrix} L_1 \\ L_2 \\ L_3 \end{bmatrix} = \begin{bmatrix} I_{11}\omega_1 + I_{12}\omega_2 + I_{13}\omega_3 \\ I_{21}\omega_1 + I_{22}\omega_2 + I_{23}\omega_3 \\ I_{31}\omega_1 + I_{32}\omega_2 + I_{33}\omega_3 \end{bmatrix} \implies L_a = I_{ab}\omega_b$$

In other words, each component L_a is a linear combination of *all* of $\boldsymbol{\omega}$'s components, which means that in general \mathbf{L} and $\boldsymbol{\omega}$ different directions. This fact leads to many of the peculiar properties of spinning objects. \mathbf{L} and $\boldsymbol{\omega}$ coincide only for rotation about a principle axis of inertia.

4.1.6 Rigid Body Motion as Translation and Rotation

Everything we've said so far applies to a body rotating about a fixed point. The fixed point of rotation has not moved through space.

- The general motion of a free rigid body can always be written as a translation through space plus a rotation about a point in space. In theory, the point of rotation could be anywhere inside or outside the body. But the rotation point is traditionally chosen to be the center of mass, because this simplifies the analysis.
- We write the position \mathbf{r}_i of a point in the body with

$$\mathbf{r}_i(t) = \mathbf{R}(t) + \Delta\mathbf{r}_i(t)$$

where $\mathbf{R}(t)$ is the center of mass and $\Delta\mathbf{r}_i(t)$ is the position of the point relative to the center of mass. With the decomposition $\mathbf{r}_i = \mathbf{R} + \Delta\mathbf{r}_i$, the kinetic energy of a rotating free body decomposes into a translation of the center of mass plus rotation about the center of mass.

$$\begin{aligned} T &= \frac{1}{2} \sum_i m_i (\dot{\mathbf{r}}_i \cdot \dot{\mathbf{r}}_i) = \frac{1}{2} \sum_i m_i [(\dot{\mathbf{R}} + \Delta\dot{\mathbf{r}}_i) \cdot (\dot{\mathbf{R}} + \Delta\dot{\mathbf{r}}_i)] \\ &= \frac{1}{2} \sum_i m_i \left(\dot{\mathbf{R}}^2 + 2\{\dot{\mathbf{R}} \cdot (\boldsymbol{\omega} \times \Delta\mathbf{r}_i)\} + (\boldsymbol{\omega} \times \Delta\mathbf{r}_i)^2 \right) \\ &= \frac{1}{2} \sum_i m_i \dot{\mathbf{R}}^2 + \frac{1}{2} \omega_a I_{ab} \omega_b \end{aligned}$$

The term in curly brackets $\{\dots\}$ is linear in $\Delta\mathbf{r}_i$ vanishes by definition of center of mass $\sum_i m_i \mathbf{r}_i = 0$.

4.1.7 Euler Equations

- Recall rigid body motion for a free body decomposes into a translation of the center of mass and a rotation about the center of mass. For now, we'll ignore the translation and focus only on rotation about the center of mass.

The angular momentum of any free body must be conserved, so $\dot{\mathbf{L}} = \mathbf{0}$, which allows us to define the body's equation of motion. Using the body-fixed expansion of angular

momentum $\mathbf{L} = L_a \mathbf{e}_a$ the equations of motion are

$$\begin{aligned} \mathbf{0} &\equiv \frac{d\mathbf{L}}{dt} = \frac{dL_a}{dt} \mathbf{e}_a + L_a \frac{d\mathbf{e}_a}{dt} \\ &= \frac{dL_a}{dt} \mathbf{e}_a + L_a \boldsymbol{\omega} \times \mathbf{e}_a \end{aligned}$$

- In principle, the body-fixed frame could be any frame rotating with the body. But the equations of motion simplify considerably if we choose the body-fixed frame to align with the system of principle axes in which \mathbf{I} is diagonal.

The $\{\mathbf{e}_a\}$ aligned with the principle axes $\mathbf{I} = \text{diag } I_1, I_2, I_3$ the angular momentum If we align the body-fixed frame with the principle axes, we get the simple expression $L_a = I_a \omega_a$ with no summation implied over a , e.g. $L_1 = I_1 \omega_1$.

The equations of motion become

$$\begin{aligned} I_1 \dot{\omega}_1 + (I_3 - I_2) \omega_2 \omega_3 &= 0 \\ I_2 \dot{\omega}_2 + (I_1 - I_3) \omega_3 \omega_1 &= 0 \\ I_3 \dot{\omega}_3 + (I_2 - I_1) \omega_1 \omega_2 &= 0 \end{aligned}$$

The equations are called the *Euler equations*. They are the equations of motion for free rigid body rotating about its center of mass; you use them to calculate the body's angular velocity $\boldsymbol{\omega}$ in the body-fixed frame. Since they really rely on \mathbf{I} being diagonal, they apply *only* if the the body-fixed frame aligns with the system of principle of axes measured about the point of rotation. If the principle axes and body-fixed frame aren't aligned, you have to use the more general equation

$$\mathbf{0} = \frac{dL_a}{dt} \mathbf{e}_a + L_a \boldsymbol{\omega} \times \mathbf{e}_a$$

Here's a quick derivation of the Euler equations using column vectors. Note that $\frac{dL_a}{dt} = \frac{d}{dt}[I_a \omega_a] = I_a \dot{\omega}_a$ because the I_a are constant in the body frame.

$$\begin{aligned} \mathbf{0} &= \frac{dL_a}{dt} \mathbf{e}_a + L_a \boldsymbol{\omega} \times \mathbf{e}_a = I_a \dot{\omega}_a \mathbf{e}_a + I_a \omega_a \boldsymbol{\omega} \times \mathbf{e}_a \\ &= \begin{bmatrix} I_1 \dot{\omega}_1 \\ I_2 \dot{\omega}_2 \\ I_3 \dot{\omega}_3 \end{bmatrix} + I_1 \omega_1 \begin{bmatrix} 0 \\ \omega_3 \\ -\omega_2 \end{bmatrix} + I_2 \omega_2 \begin{bmatrix} -\omega_3 \\ 0 \\ \omega_1 \end{bmatrix} + I_3 \omega_3 \begin{bmatrix} \omega_2 \\ -\omega_1 \\ 0 \end{bmatrix} \\ &= \begin{bmatrix} I_1 \dot{\omega}_1 + (I_3 - I_2) \omega_2 \omega_3 \\ I_2 \dot{\omega}_2 + (I_1 - I_3) \omega_3 \omega_1 \\ I_3 \dot{\omega}_3 + (I_2 - I_1) \omega_1 \omega_2 \end{bmatrix} = \begin{bmatrix} 0 \\ 0 \\ 0 \end{bmatrix} \end{aligned}$$

- We can also generalize Euler's equations to apply to a body acted on by an external torque \mathbf{M} . Instead of starting with $\dot{\mathbf{L}} = \mathbf{0}$ we start from $\dot{\mathbf{L}} = \mathbf{M}$, expanding the vectors in the body-fixed frame and following an analogous derivation. The Euler equations with an external torque are

$$\begin{aligned} I_1 \dot{\omega}_1 + (I_3 - I_2) \omega_2 \omega_3 &= M_1 \\ I_2 \dot{\omega}_2 + (I_1 - I_3) \omega_3 \omega_1 &= M_2 \\ I_3 \dot{\omega}_3 + (I_2 - I_1) \omega_1 \omega_2 &= M_3 \end{aligned}$$

where M_1, M_2 and M_3 are the components of the torque in the body-fixed frame.

- I stress again that the Euler equations hold *only* if the body-fixed frame aligns with the principle axes. The most general equation of motion for rigid body rotation is the vector equation

$$\mathbf{M} = \dot{\mathbf{L}} = \frac{d[\mathbf{I}\boldsymbol{\omega}]}{dt} = \mathbf{I}\dot{\boldsymbol{\omega}} + \boldsymbol{\omega} \times (\mathbf{I}\boldsymbol{\omega})$$

where \mathbf{I} is the body's inertia tensor about the point of rotation, \mathbf{M} is the externally applied torque and $\boldsymbol{\omega}$ is the body's angular velocity. Of course, all three quantities \mathbf{I} , $\boldsymbol{\omega}$ and \mathbf{M} must be measured with respect to the same reference point. In the body-fixed system of principle axes, this equation simplifies to the Euler equations.

4.2 Motion of the Free Top

A free top is fancy name for a freely rotating rigid body with no external torques acting on it.

4.2.1 Symmetric Free Top

- A *symmetric top* is a rotating rigid body with two equal principle moments are equal to each other e.g. $I_1 = I_2$, while the third moment I_3 is different from I_1 and I_2 . The \mathbf{e}_3 axis is the top's axis of symmetry, and is sometimes called the *figure axis*.

For a free symmetric top with $I_1 = I_2$, the Euler equations become

$$\begin{aligned} I_1 \dot{\omega}_1 &= \omega_2 \omega_3 (I_1 - I_3) \\ I_1 \dot{\omega}_2 &= -\omega_3 \omega_1 (I_1 - I_3) \\ I_3 \dot{\omega}_3 &= 0 \end{aligned}$$

ω_3 is the component (measured in the body-fixed system of principle axis) of the angular velocity about the axis of symmetry \mathbf{e}_3 and, from the third Euler equation $I_3 \dot{\omega}_3 = 0$, is constant for a free symmetric top.

The components ω_1 and ω_2 are given by the coupled first-order differential equations

$$\dot{\omega}_1 = \Omega \omega_2 \quad \text{and} \quad \dot{\omega}_2 = -\Omega \omega_1 \quad \text{where } \Omega = \frac{(I_1 - I_3)}{I_1} \omega_3$$

The equations are solved by

$$\omega_1 = \omega_0 \sin \Omega t \quad \text{and} \quad \omega_2 = \omega_0 \cos \Omega t$$

This solution has a nice physical interpretation: in the body-fixed frame, the free top's angular velocity vector $\boldsymbol{\omega}$ precesses about the figure axis \mathbf{e}_3 with angular frequency Ω . The direction of $\boldsymbol{\omega}$'s rotation depends on the sign of Ω and thus on whether $I_1 > I_3$ or $I_3 > I_1$. The precession of $\boldsymbol{\omega}$ about the figure axis \mathbf{e}_3 in the body-fixed frame is sometimes called *wobble*.

- What would we see in the lab frame? We know the angular momentum \mathbf{L} is constant since the top is free. More so, ω_3 is constant, so $L_3 = I_3 \omega_3$ is also constant. Since both \mathbf{L} and L_3 are constant, the angle between the figure axis \mathbf{e}_3 and angular momentum \mathbf{L} is constant. The figure axis \mathbf{e}_3 thus precesses about the fixed angular momentum \mathbf{L} , while the body rotates so that $\boldsymbol{\omega}$ remains between \mathbf{e}_3 and \mathbf{L} . [Click here for a nice animation.](#)

4.2.2 Stability and Perturbations of The Asymmetric Top

- An asymmetric top basically means a rigid body without any symmetries. This means none of the principal moments of inertia are equal, i.e. $I_1 \neq I_2 \neq I_3 \neq I_1$. The rotational motion of an asymmetric top is much more complicated than a free top's.
- We will analyze only a simple case: when the a free asymmetric top rotating about a principle axis of inertia, e.g. \mathbf{e}_1 so $\boldsymbol{\omega}$ and \mathbf{e}_1 have the same direction. In this case, in the body-fixed system of principle axes

$$\omega_1 = \Omega \quad \text{and} \quad \omega_2 = \omega_3$$

This solves Euler's equations.

- We then ask what happens if the top's angular velocity is displaced slightly from the principle axis \mathbf{e}_1 , considering small perturbations about the angular velocity of the form

$$\omega_1 = \Omega + \eta_1 \quad \omega_2 = \eta_2 \quad \omega_3 = \eta_3$$

where the $\eta_a = \eta_a(t)$ and $\eta_a \ll \Omega$. We then substitute the ω_a into the Euler equations. Since the η_a are small, we ignore terms of order η_2 and higher, e.g. $\eta_2\eta_3 \approx 0$ in the first equation. The result is

$$\begin{aligned} I_1 \dot{\eta}_1 &= 0 \\ I_2 \dot{\eta}_2 &= \Omega \eta_3 (I_1 - I_3) \\ I_3 \dot{\eta}_3 &= \Omega \eta_2 (I_2 - I_1) \end{aligned}$$

Differentiating the second equation and solving for $\dot{\eta}_3$, then substituting $\dot{\eta}_3$ into the third equation gives

$$I_2 \ddot{\eta}_2 = \frac{\Omega^2}{I_3} (I_1 - I_2)(I_3 - I_1) \eta_2 \equiv A \eta_2$$

Which has the familiar form $\ddot{x} = kx$. The stability of the perturbation depends on the sign of A :

- If $A < 0$ the disturbance will oscillate around the position of constant motion.
- If $A > 0$ the disturbance will grow exponentially, and the system is unstable.

From the definition $A \equiv \frac{\Omega^2}{I_3} (I_1 - I_2)(I_3 - I_1)$, the system is unstable ($A > 0$) if

$$I_2 < I_1 < I_3 \quad \text{or} \quad I_3 < I_1 < I_2$$

while all other possibilities are stable. Physically, this means a body will rotate stably about the principle axis with the largest or the smallest moment of inertia, but not about the intermediate axis.

4.3 Euler Angles

- Informally, Euler angles are just coordinates well-suited to describing the dynamics of a rigid body.

More exactly, the Euler angles are a way of parametrizing the rotation \mathbf{R} from the inertial lab frame $\{\tilde{\mathbf{e}}_a\}$ to the body-fixed frame $\{\mathbf{e}_a\}$, which is discussed at the beginning of the chapter. The angles describe how the body-fixed system of principal axes of a rigid body is oriented relative to an inertial lab frame.

The Euler angles rest on theoretical framework of *Euler's theorem*, which tells us:

An arbitrary rotation in three dimensions can be expressed as the product of 3 successive rotations about 3 (in general) different axes.

- Calling on Euler's theorem, we can get from the inertial lab system $\{\tilde{\mathbf{e}}_a\}$ to the body-fixed system $\{\mathbf{e}_a\}$ in three steps, using the rotations $\mathbf{R}_3(\phi), \mathbf{R}_1(\theta)$ and $\mathbf{R}_3(\psi)$

$$\{\tilde{\mathbf{e}}_a\} \xrightarrow{\mathbf{R}_3(\phi)} \{\mathbf{e}'_a\} \xrightarrow{\mathbf{R}_1(\theta)} \{\mathbf{e}''_a\} \xrightarrow{\mathbf{R}_3(\psi)} \{\mathbf{e}_a\}$$

where $\mathbf{R}_a(\phi)$ means a rotation about the a th basis vector by the angle ϕ , e.g. $\mathbf{R}_3(\phi)$ means a rotation about $\tilde{\mathbf{e}}_3$ by the angle ϕ . The angles of rotation ϕ, θ and ψ are called the *Euler angles*. The three steps in explained in more detail below:

1. Rotate about $\tilde{\mathbf{e}}_3$ by the angle ϕ . ϕ is called the precession angle. We have

$$\mathbf{e}'_a = \mathbf{R}_3(\phi)_{ab} \tilde{\mathbf{e}}_b \quad \mathbf{R}_3(\phi) = \begin{bmatrix} \cos \phi & \sin \phi & 0 \\ -\sin \phi & \cos \phi & 0 \\ 0 & 0 & 1 \end{bmatrix}$$

2. Rotate about \mathbf{e}'_1 by the angle θ . θ is called the nutation angle.

$$\mathbf{e}''_a = \mathbf{R}_1(\theta)_{ab} \mathbf{e}'_b \quad \mathbf{R}_1(\theta) = \begin{bmatrix} 1 & 0 & 0 \\ 0 & \cos \theta & \sin \theta \\ 0 & -\sin \theta & \cos \theta \end{bmatrix}$$

3. Rotate about \mathbf{e}''_3 by angle ψ to reach the body-fixed frame $\{\mathbf{e}_a\}$. ψ corresponds to rotation of the body about its figure axis.

$$\mathbf{e}_a = \mathbf{R}_3(\psi)_{ab} \mathbf{e}''_b \quad \mathbf{R}_3(\psi) = \begin{bmatrix} \cos \psi & \sin \psi & 0 \\ -\sin \psi & \cos \psi & 0 \\ 0 & 0 & 1 \end{bmatrix}$$

- The angles ϕ, θ, ψ are the Euler angles. The complete rotation matrix transforming from the lab frame to the body-fixed frame is

$$\mathbf{R}(\phi, \theta, \psi) = \mathbf{R}_3(\psi) \mathbf{R}_1(\theta) \mathbf{R}_3(\phi) \quad \text{and} \quad \mathbf{e}_a = \mathbf{R}_{ab}(\phi, \theta, \psi) \tilde{\mathbf{e}}_b$$

In full, the matrix reads

$$\mathbf{R}(\phi, \theta, \psi) = \begin{bmatrix} \cos \psi \cos \phi - \cos \theta \sin \phi \sin \psi & \sin \phi \cos \psi + \cos \theta \sin \psi \cos \phi & \sin \theta \sin \psi \\ -\cos \phi \sin \psi - \cos \theta \cos \psi \sin \phi & -\sin \psi \sin \phi + \cos \theta \cos \psi \cos \phi & \sin \theta \cos \psi \\ \sin \theta \sin \phi & -\sin \theta \cos \phi & \cos \theta \end{bmatrix}$$

- To expand a vector \mathbf{r} in the body frame: $\mathbf{r} = r_a \mathbf{e}_a$.

To expand \mathbf{r} in the lab frame: $\mathbf{r} = \tilde{r}_a \tilde{\mathbf{e}}_a$ where $\tilde{r}_b = r_a \mathbf{R}(\phi, \theta, \psi)_{ab}$

4.3.1 Angular Velocity and Kinetic Energy

The goal is to write $\boldsymbol{\omega}$ in the body fixed system in the form $\boldsymbol{\omega} = w_1 \mathbf{e}_1 + w_2 \mathbf{e}_2 + w_3 \mathbf{e}_3$ where $\omega_{x,y,z} = \omega_{x,y,z}(\phi, \theta, \psi)$ are written in terms of the Euler angles.

- By the construction of the Euler angles, the angular velocity can be written

$$\boldsymbol{\omega} = \dot{\phi}\tilde{\mathbf{e}}_3 + \dot{\theta}\mathbf{e}'_1 + \dot{\psi}\mathbf{e}_3$$

where the basis vectors are the axes of rotation in each step of the Euler angles: $\tilde{\mathbf{e}}_3$ is the lab-frame z axis, \mathbf{e}'_1 is the intermediate x' axis after the first rotation, and \mathbf{e}_3 is the object's figure axis.

- In terms of the body-fixed frame $\{\mathbf{e}_a\}$, the basis vectors $\tilde{\mathbf{e}}_3, \mathbf{e}'_1$ and \mathbf{e}_3 read

$$\tilde{\mathbf{e}}_3 = \sin\theta \sin\psi \mathbf{e}_1 + \sin\theta \cos\psi \mathbf{e}_2 + \cos\theta \mathbf{e}_3$$

$$\mathbf{e}'_1 = \cos\psi \mathbf{e}_1 - \sin\psi \mathbf{e}_2$$

$$\mathbf{e}_3 = \mathbf{e}_3$$

- Plugging the basis vectors into $\boldsymbol{\omega} = \dot{\phi}\tilde{\mathbf{e}}_3 + \dot{\theta}\mathbf{e}'_1 + \dot{\psi}\mathbf{e}_3$ and grouping like terms produces

$$\boldsymbol{\omega} = (\dot{\phi} \sin\theta \sin\psi + \dot{\theta} \cos\psi) \mathbf{e}_1 + (\dot{\phi} \sin\theta \cos\psi - \dot{\theta} \sin\psi) \mathbf{e}_2 + (\dot{\phi} \cos\theta + \dot{\psi}) \mathbf{e}_3$$

the desired expression for angular velocity in the body-fixed system.

Kinetic Energy: In terms of Euler angles *in a body-fixed system in which the inertia tensor is diagonal*, the body's kinetic energy reads

$$T = \frac{1}{2} (I_x \omega_x^2 + I_y \omega_y^2 + I_z \omega_z^2)$$

where $\omega_x, \omega_y, \omega_z$ are given in terms of the Euler angles as above. If additionally the top is symmetric so that $\mathbf{I} = (I_1, I_1, I_3)$, the kinetic energy simplifies to

$$T = \frac{I_1}{2} (\dot{\phi}^2 \sin^2\theta + \dot{\theta}^2) + \frac{I_3}{2} (\dot{\phi} \cos\theta + \dot{\psi})^2$$

4.3.2 Free Symmetric Top In Terms of Euler Angles

Earlier we analyzed the free symmetric top in the body-fixed frame $\{\mathbf{e}_a\}$. This is take two, using the Euler angles and working in the *lab* frame $\{\tilde{\mathbf{e}}_a\}$.

- We choose the angular momentum \mathbf{L} to align with the $\tilde{\mathbf{e}}_3$ axis. The top is free, so \mathbf{L} is conserved, as is the body-fixed frame component $L_3 = I_3 \omega_3$; since both quantities are conserved, the angle between them is also conserved.

This angle is precisely the second Euler angle, i.e. the nutation angle θ , so $\dot{\theta} = 0$.

A free symmetric top does not perform nutation when \mathbf{L} coincides with $\tilde{\mathbf{e}}_3$.

- The next step is to use the Euler-angle expression for angular velocity to solve for $\dot{\psi}$; the result is $\dot{\psi} = \Omega \equiv \frac{(I_1 - I_3)}{I_1} \omega_3$, meaning the object rotates about its figure axis with angular frequency Ω .
- Finally, using $\dot{\psi} = \Omega$ in the expression $\omega_3 = \dot{\phi} \cos\theta + \dot{\psi}$, we solve for the precession frequency $\dot{\phi}$. The result is

$$\dot{\phi} = \frac{\omega_3 - \Omega}{\cos\theta} = \frac{I_3 \omega_3}{I_1 \cos\theta}$$

4.4 Heavy Symmetric Top

A heavy top is a top acted on by gravity. Symmetric means that two of the principle moments of inertia are equal, so in the body-fixed system of principle axes we have $\mathbf{I} = (I_1, I_1, I_3)$. This analysis assumes the top is fixed at some point (e.g. on the ground) a distance l from the center of mass.

4.4.1 Lagrangian and Conserved Quantities

- The kinetic energy is $T = \frac{1}{2}I_1(\omega_1 + \omega_2)^2 + \frac{1}{2}I_3\omega_3^2$ and the potential energy is $V = mgl \cos \theta$. In terms of the Euler angles, the heavy symmetric top's Lagrangian is

$$\begin{aligned} L &= \frac{1}{2}I_1(\dot{\phi}^2 + \dot{\psi}^2) + \frac{1}{2}I_3\omega_3^2 - mgl \cos \theta \\ &= \frac{1}{2}I_1(\dot{\phi}^2 \sin^2 \theta + \dot{\theta}^2) + \frac{1}{2}I_3(\dot{\phi} \cos \theta + \dot{\psi})^2 - mgl \cos \theta \end{aligned}$$

- Both ϕ and ψ are cyclic coordinates and lead to the conserved quantities

$$\begin{aligned} p_\phi &= I_1 \sin^2 \theta \dot{\phi} + I_3 \cos \theta (\dot{\psi} + \dot{\phi} \cos \theta) \\ p_\psi &= I_3 (\dot{\psi} + \cos \theta \dot{\phi}) = I_3 \omega_3 \end{aligned}$$

p_ψ is the top's angular momentum about the figure axis \mathbf{e}_3 . Like for the free symmetric top, the angular velocity component ω_3 about the figure axis is constant. This makes sense; there is no external torque acting on the top about its figure axis.

Energy

- The top's total energy $E = T + V$ is also conserved, as for any isolated mechanical system. For future use, we define the constants

$$a = \frac{I_3}{I_1} \omega_3 \quad b = \frac{p_\phi}{I_1}$$

In terms of a and b , the derivatives $\dot{\phi}$ and $\dot{\psi}$ are

$$\dot{\phi} = \frac{b - a \cos \theta}{\sin^2 \theta} \quad \dot{\psi} = \frac{I_1 a}{I_3} - \frac{(b - a \cos \theta) \cos \theta}{\sin^2 \theta}$$

If $\theta = \theta(t)$ were known, we could theoretically integrate these two expressions to solve for $\phi(t)$ and $\psi(t)$. So the next step is to find $\theta(t)$.

Reduced Energy

- We define a reduced energy E' via $E = E' + \frac{1}{2}I_3\omega_3^2$. Since E , I_3 and ω_3 are constant, E' is also constant.

$$E' = \frac{1}{2}I_1\dot{\theta}^2 + V_{\text{eff}}(\theta), \quad V_{\text{eff}}(\theta) = \frac{I_1(b - a \cos \theta)^2}{2 \sin^2 \theta} + mgl \cos \theta$$

We then introduce a new coordinate $u = \cos \theta$, where $u \in [-1, 1]$ because $\theta \in [0, \pi]$, and two more constants

$$\alpha = \frac{2E'}{I_1} \quad \text{and} \quad \beta = \frac{2mgl}{I_1}$$

- The equations of motion for $u = \cos \theta$, ϕ , ψ can now be written

$$\begin{aligned} \dot{u}^2 &= (1 - u^2)(\alpha - \beta u) - (b - au)^2 \equiv f(u) \\ \dot{\phi} &= \frac{b - au}{1 - u^2} \\ \dot{\psi} &= \frac{I_1 a}{I_3} - \frac{u(b - au)}{1 - u^2} \end{aligned}$$

In theory, the equation for \dot{u}^2 could be solved analytically in terms of elliptic integrals, and then the time dependence of θ could be implicitly extracted using $u = \cos \theta$. With θ in hand, we could then solve for ϕ and ψ .

- But there's a better, more physical, approach: using the function $f(u)$ as a tool to analyze the nutation angle θ . Notice that $\dot{u}^2 = f(u)$ behaves as a cubic polynomial in u . A quick look at the limits $u \rightarrow \pm\infty$ shows

$$f(u) \rightarrow \begin{cases} +\infty & \text{as } u \rightarrow \infty \\ -\infty & \text{as } u \rightarrow -\infty \end{cases}$$

Also important is the behavior at ± 1 :

$$f(\pm 1) = -(b \mp a)^2 \leq 0$$

Although $f(u)$ is mathematically defined on the entire real line, the region of *physical* interest is, happily, much smaller. Values of u are constrained to positive values of $f(u) = \dot{u}^2$ (corresponding to real values of \dot{u}) in the region $u \in [-1, 1]$ (because $|u| = |\cos \theta| \leq 1$).

We'll study three special cases of motion in the next three subsections.

4.4.2 Letting the Top Go

- In this case, sometime called the “released top”, we spin the top about its axis of symmetry and let it go at some angle $\theta \neq 0$ with $\dot{\theta} = 0$. The initial conditions are

$$\begin{aligned} \dot{\theta}(0) = 0 &\implies \dot{u}_{t=0}^2 = 0 \implies f(u_{t=0}) = 0 \\ \dot{\phi}(0) = 0 &\implies b - au_{t=0} = 0 \implies u_{t=0} = \frac{b}{a} \end{aligned}$$

- Recall the quantity $p_\phi = I_1 \dot{\phi} \sin^2 \theta + I_3 \omega_3 \cos \theta$ is a constant of motion. Evaluating p_ϕ at $t = 0$ when $\dot{\phi} = 0$ and $\theta = \theta_0$ gives the relationship

$$p_\phi = I_1 \dot{\phi} \sin^2 \theta + I_3 \omega_3 \cos \theta = I_3 \omega_3 \cos \theta_0$$

This is enough to determine the qualitative motion of the top. When released, it starts to fall under the influence of gravity, so θ increases. As the top falls and θ increases (and thus $I_3 \omega_3 \cos \theta$ decreases) the term $I_1 \dot{\phi} \sin^2 \theta$ must increase to keep p_ϕ constant, resulting as an increase in $\dot{\phi}$.

More so, to keep the signs of $\dot{\phi}$ and ω_3 balanced, the direction of precession must in the same as the direction of spin about the figure axis.

4.4.3 Uniform Precession Without Nutation

- This means the top spins uniformly without bobbing up and down. Mathematically, this means $\dot{\phi}$ is constant (uniform precession) and $\dot{\theta} = 0$ (no nutation).

Recall the relationship $\dot{\theta}^2 = f(u)$. To satisfy $\dot{\theta} = 0$ for all time, we need $f(u)$ to have only a single root u_0 and no other values of $f(u)$ in the physical range $u \in [-1, 1]$ and $u \geq 0$ which constrains the top's physical motion to a single point of constant $\dot{\theta} = 0$. The root must satisfy

$$\begin{aligned} f(u_0) &= (1 - u_0^2)(\alpha - \beta u_0) - (b - au_0)^2 = 0 \\ f'(u_0) &= -2u_0(\alpha - \beta u_0) - \beta(1 - u_0^2) + 2a(b - au_0) = 0 \end{aligned}$$

Combining the equations gives the condition $\frac{1}{2}\beta = a\dot{\phi} - \dot{\phi}^2 u_0$, and substituting in the definitions $I_1 a = I_3 \omega_3$ and $\beta = \frac{2Mgl}{I_1}$ gives

$$Mgl = \dot{\phi}(I_3 \omega_3 - I_1 \dot{\phi} \cos \theta_0)$$

- Here is the physical interpretation: for a fixed value of ω_3 (the top's spin about its figure axis) and θ_0 (the angle from the vertical at which its released) there is an exact value of $\dot{\phi}$ (a “push” in the direction of precession) for which the top will spin without nutation, i.e. satisfying $Mgl = \dot{\phi}(I_3 \omega_3 - I_1 \dot{\phi} \cos \theta_0)$.
- Because the condition $Mgl = \dot{\phi}(I_3 \omega_3 - I_1 \dot{\phi} \cos \theta_0)$ is quadratic in $\dot{\phi}$, there are in principle two values of $\dot{\phi}$ for which the top could precess without nutation.

If you rearrange and calculate the discriminant, you'll see equation has real solutions for $\dot{\phi}$ if

$$\omega_3 > \frac{2}{I_3} \sqrt{Mgl I_1 \cos \theta_0}$$

In other words ω_3 must be large enough (the top must be spinning fast enough about its figure axis) to be able to precess without nutation at all. What happens when the top isn't spinning fast enough? It falls over!

4.4.4 The Sleeping Top

- In this case, we start the top spinning in an upright position with $\theta = \dot{\theta} = 0$. When the top spins completely upright with $\theta = 0$, it is called a *sleeping top*. But does the top stay there?

Turning to the definitions $\dot{u}^2 = f(u)$ and $u = \cos \theta$, the sleeping top conditions $\theta = \dot{\theta} = 0$ are met if $u = 1$ (for $\theta = 0$) and $f(1) = 0$ (for $\dot{\theta} = 0$).

In these conditions $a = b$ and $\alpha = \beta$. Here's a quick confirmation; feel free to skip.

$$\begin{aligned} a|_{\theta=0} &= \frac{I_3 \omega_3|_{\theta=0}}{I_1} = \frac{I_3(\dot{\psi} + \dot{\phi})}{I_1} = \frac{p_\phi|_{\theta=0}}{I_1} = b|_{\theta=0} \\ \alpha|_{\theta=0} &= \frac{2E'|_{\theta=0}}{I_1} = \frac{2(0 + Mgl)}{I_1} = \beta \end{aligned}$$

- With $a = b$ and $\alpha = \beta$ the function $f(u)$ simplifies to

$$f(u) = (1 - u)^2 [\alpha(1 - u) - a^2]$$

so $f(u)$ has a double zero at $u = +1$, while a second zero occurs at $u_2 = \frac{a^2}{\alpha - 1}$. There are two possibilities:

- If $u_2 > 1$ (i.e. outside the region of physical motion) and $\omega_3^2 > \frac{4I_1 Mgl}{I_3^2}$ the motion is stable: the only allowed physically allowed value of u corresponds to $\theta = 0$, since all other values of $f(u)$ are either negative or occur outside of $[-1, 1]$.

For slight perturbations from the initial conditions, the nutation will stably oscillate with small displacements about $\theta = 0$.

- If $u_2 < 1$ (inside the region of physical motion) and $\omega_3^2 < \frac{4I_1Mgl}{I_3^2}$, there is a region of allowed physical motion to the left of the zero $u = 1$. The motion is unstable: if the top is displaced slightly from $\theta = 0$, it will fall off into the region of decreasing u and thus increasing θ .

Notice that the top is stable for fast spins ω_3 about the figure axis and unstable for slow spins. In practice, a sleeping top stays upright until ω_3 is slowed by friction to

$$\omega_3 = \frac{4I_1Mgl}{I_3^2}$$

at which point the top starts to fall and precess.

5 Hamiltonian Mechanics

5.1 Hamilton's Equations

- The foundation of the Lagrangian formalism is the Lagrangian function $L(q_i, \dot{q}_i, t)$ where $q_i, i = 1, \dots, n$ are the n generalized coordinates describing our system. The equations of motion are

$$\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) - \frac{\partial L}{\partial q_i} = 0$$

In the Lagrangian formalism the equations of motion are n second order differential equations; for n second order equations we need $2n$ initial conditions to define our system; typically there are $q_i(0)$ and $\dot{q}_i(0)$.

- The Hamiltonian formalism also uses n generalized coordinates q_i . However, instead of the derivatives \dot{q}_i , Hamiltonian mechanics uses the *generalized momenta*

$$p_i \equiv \frac{\partial L}{\partial \dot{q}_i}$$

Because they are derived from the Lagrangian, the generalized momenta are also functions of q_i, \dot{q}_i and t ; we have $p_i = p_i(q_j, \dot{q}_j, t)$.

- In terms of the generalized momenta $p_i = \frac{\partial L}{\partial \dot{q}_i}$ the Lagrangian equations of motion read

$$\frac{d}{dt} - \frac{\partial L}{\partial q_i} = 0 \implies \dot{p}_i = \frac{\partial L}{\partial q_i}$$

Notice we have eliminated the derivatives \dot{q}_i in favor of the momenta p_i .

- A conceptual approach to what's going on: recall the set $\{q_i\}$ defines a point in our system's n -dimensional configuration space C , and the time evolution of our system is represented by an n -dimensional path in C .

However, the *state* of the system is defined by both $\{q_i\}$ and $\{p_i\}$ in the sense that knowing both the coordinates and momenta allows us to determine the system's state at all future times. The pair $\{q_i, p_i\}$ represents a point in our system's $2n$ -dimensional *phase space*. Knowing a single point in phase space uniquely determines the system's evolution in time; since a point *uniquely* defines the system's future, paths in phase space never cross.

5.1.1 Mathematical Detour: The Legendre Transform

- We want to find a function defined on a system's phase space that will determine the unique evolution of q_i and p_i . Such a function must naturally be a function of q_i and p_i , but must contain the same information as the Lagrangian $L(q_i, \dot{q}_i, t)$ (which we know also determines the system's time evolution). We create our desired function with a mathematical technique called the *Legendre transform*. The Legendre transform goes something as follows:
- First, consider an arbitrary function $f(x, y)$, so the total derivative is

$$df = \frac{\partial f}{\partial x} dx + \frac{\partial f}{\partial y} dy$$

Next we define a new function $g(x, y, u) = ux - f(x, y)$ of three variables x, y and u . The total derivative dg is

$$dg = d[ux] - df = u dx + x du - \frac{\partial f}{\partial x} dx - \frac{\partial f}{\partial y} dy$$

- The total derivative simplifies considerably for a special choice of u . It is $u(x, y) = \frac{\partial f}{\partial x}$. In this case the terms $u dx$ and $\frac{\partial f}{\partial x} dx$ vanish. We have

$$dg = x du - \frac{\partial f}{\partial y} dy$$

Notice that all instances of dx are gone; only du and dy remain. In other words, g is naturally thought of as a function of only u and y , i.e. $g = g(u, y)$.

- To get an explicit expression for $g(u, y)$, we first solve for x by inverting $u(x, y) = \frac{\partial f}{\partial x}$ to get $x = x(u, y)$, then insert this into the definition $g = ux - f$ to get

$$g(u, y) = ux(u, y) - f(x(u, y), y)$$

This last result is the Legendre transform. It takes us from the function $f(x, y)$ to the function $g(u, y)$, eliminating the dependence on x with $u(x, y) = \frac{\partial f}{\partial x}$.

- Importantly, the Legendre transform g preserves all information about the original f . Concretely, we can recover $f(x, y)$ from $g(u, y)$ and $u(x, y) = \frac{\partial f}{\partial x}$ with

$$\left. \frac{\partial g}{\partial u} \right|_y = x(u, y) \quad \text{and} \quad \left. \frac{\partial g}{\partial y} \right|_u = -\frac{\partial f}{\partial y}$$

which means the inverse Legendre transform

$$f(x, y) = \left(\frac{\partial g}{\partial u} \right) u - g(u, y)$$

recovers the original function f .

5.1.2 Hamilton's Equations of Motion

- The Lagrangian $L(q_i, \dot{q}_i, t)$ was a function of the system's coordinates, their time derivatives, and possible time. In Hamiltonian mechanics, the Lagrangian is replaced by the *Hamiltonian*, which is a Legendre transform of L to eliminate \dot{q}_i defined as

$$H(q_i, p_i, t) = \sum_{i=1}^n p_i \dot{q}_i - L(q_i, \dot{q}_i, t)$$

- Where do we see the elimination of q_i ? Returning to the Einstein summation notation we have

$$dH = p_i d\dot{q}_i + \dot{q}_i dp_i - \left(\frac{\partial L}{\partial q_i} dq_i + \frac{\partial L}{\partial \dot{q}_i} d\dot{q}_i + \frac{\partial L}{\partial t} dt \right)$$

From the construction $p_i = \frac{\partial L}{\partial \dot{q}_i}$, the terms $p_i d\dot{q}_i$ and $-\frac{\partial L}{\partial \dot{q}_i} d\dot{q}_i$ cancel. We have

$$dH = \dot{q}_i dp_i - \frac{\partial L}{\partial q_i} dq_i - \frac{\partial L}{\partial t} dt$$

and indeed we see $H = H(q_i, p_i, t)$. More so, we can compare our expression for dH to the general total differential $dH = \frac{\partial H}{\partial q_i} dq_i + \frac{\partial H}{\partial p_i} dp_i + \frac{\partial H}{\partial t} dt$ to conclude

$$\frac{\partial H}{\partial p_i} = \dot{q}_i \quad \frac{\partial H}{\partial t} = -\frac{\partial L}{\partial t} \quad \frac{\partial H}{\partial q_i} = -\frac{\partial L}{\partial q_i} = -\frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right) \equiv -\dot{p}_i$$

where we use the Euler-Lagrange equations $\frac{\partial L}{\partial q_i} = \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}_i} \right)$ in the last equality.

- The last results are important: they are the equations of motion in the Hamiltonian formalism. They are most commonly written

$$\dot{q}_i = \frac{\partial H}{\partial p_i} \quad \dot{p}_i = -\frac{\partial H}{\partial q_i}$$

Note that the Hamiltonian formalism replaces the Lagrangian's n second order differential equations with $2n$ first order equations for q_i and p_i .

5.1.3 Examples Using the Hamiltonian Equations of Motion

A Particle in a Potential Field

- The generalized coordinate is simply the position vector $\mathbf{r} = (x, y, z)$ and the Lagrangian L is

$$L = \frac{1}{2} m \dot{\mathbf{r}}^2 - V(\mathbf{r})$$

- The conjugate momenta $p = \frac{\partial L}{\partial \dot{q}}$ are

$$\mathbf{p} = m \dot{\mathbf{r}}$$

which in this case coincides with the usual definition of momentum.

- The Hamiltonian $H = p_i \dot{q}_i - L$ is

$$H = \mathbf{p} \cdot \dot{\mathbf{r}} - \left(\frac{1}{2} m \dot{\mathbf{r}}^2 - V(\mathbf{r}) \right) = \frac{\mathbf{p}^2}{2m} + V(\mathbf{r})$$

Notice the Hamiltonian is written as a function of \mathbf{p} in favor of $\dot{\mathbf{q}}$.

- The Hamiltonian equations of motion are

$$\dot{\mathbf{r}} = \frac{\partial H}{\partial \mathbf{p}} = \frac{\mathbf{p}}{m} \quad \text{and} \quad \dot{\mathbf{p}} = -\frac{\partial H}{\partial \mathbf{r}} = -\nabla V$$

These equations are familiar; the first is the definition of momentum $\mathbf{p} = m \dot{\mathbf{r}}$ while the second is Newton's second law of motion for a particle in a potential field.

A Charged Particle in an Electromagnetic Field

Here is the charged particle in an EM field à la Hamilton. We studied the Lagrangian approach in Subsection 2.6

- Recall that electric field \mathbf{E} and magnetic field \mathbf{B} can be written in terms of a scalar potential $\phi(\mathbf{r}, t)$ and a vector potential $\mathbf{A}(\mathbf{r}, t)$

$$\mathbf{E} = -\nabla \phi - \frac{\partial \mathbf{A}}{\partial t} \quad \text{and} \quad \mathbf{B} = \nabla \times \mathbf{A}$$

The Lagrangian for a particle of charge e in an EM field is

$$L = \frac{1}{2}m\dot{\mathbf{r}}^2 - e(\phi - \dot{\mathbf{r}} \cdot \mathbf{A})$$

while the momentum \mathbf{p} conjugate to the position \mathbf{r} is

$$\mathbf{p} = \frac{\partial L}{\partial \dot{\mathbf{r}}} = m\dot{\mathbf{r}} + e\mathbf{A}$$

- Inverting the expression for \mathbf{p} gives us

$$\dot{\mathbf{r}} = \frac{1}{m}(\mathbf{p} - e\mathbf{A})$$

from which we can calculate the Hamiltonian

$$\begin{aligned} H(\mathbf{p}, \mathbf{r}) &= \mathbf{p}\dot{\mathbf{r}} - L = \frac{\mathbf{p}}{m}(\mathbf{p} - e\mathbf{A}) - \left[\frac{1}{2m}(\mathbf{p} - e\mathbf{A})^2 - e\phi + \frac{e}{m}(\mathbf{p} - e\mathbf{A}) \cdot \mathbf{A} \right] \\ &= \frac{1}{2m}(\mathbf{p}^2 - 2e\mathbf{p} \cdot \mathbf{A} + e^2\mathbf{A}^2) + e\phi \\ &= \frac{1}{2m}(\mathbf{p} - e\mathbf{A})^2 + e\phi \end{aligned}$$

- The first equation of motion is

$$\dot{\mathbf{r}} = \frac{\partial H}{\partial \mathbf{p}} = \frac{1}{m}(\mathbf{p} - e\mathbf{A})$$

The equation for $\dot{\mathbf{p}} = -\frac{\partial H}{\partial \mathbf{r}}$ is best given in terms of indices and components

$$\dot{p}_a = -\frac{\partial H}{\partial r_a} = \frac{e}{m}(p_b - eA_b)\frac{\partial A_b}{\partial r_a} - e\frac{\partial \phi}{\partial r_a}$$

some manipulation of the indices and the definition of the cross product in component form shows this is equivalent to the Lorentz force $\mathbf{F} = e(\mathbf{E} - \mathbf{v} \times \mathbf{B})$.

We now consider a concrete example of a particle in a uniform one-directional magnetic field $\mathbf{B} = (0, 0, B)$.

- The magnetic field $\mathbf{B} = (0, 0, B)$ corresponds to a vector potential

$$\mathbf{A} = (-By, 0, 0)$$

Consider a particle moving in the (x, y) plane. Then $\mathbf{p} = (p_x, p_y, 0)$ and the Hamiltonian is

$$\begin{aligned} H &= \frac{1}{2m} [p_x^2 + p_y^2 + 2ep_xBy + e^2B^2y^2] \\ &= \frac{1}{2m}(p_x + eBy)^2 + \frac{p_y^2}{2m} \end{aligned}$$

- The equations of motion are

$$\begin{aligned} \dot{x} &= \frac{\partial H}{\partial p_x} = \frac{1}{m}(p_x + eBy) & \dot{y} &= \frac{\partial H}{\partial p_y} = \frac{p_y}{m} & \dot{z} &= 0 \\ \dot{p}_y &= -\frac{\partial H}{\partial y} = -\frac{eB}{m}(p_x + eBy) & \dot{p}_x &= \dot{p}_z = 0 \end{aligned}$$

- Next, we note that

$$eBx + m\dot{y} = \text{constant} \equiv a \quad \text{and} \quad p_x = m\dot{x} - eBy = \text{constant} \equiv b$$

The first equality can be found by combining the equations for \dot{x} and \dot{p}_y , integrating and using $p_y = m\dot{y}$; the second e.g. by definition $p_x = \frac{\partial L}{\partial \dot{x}}$ from the Lagrangian.

- The result is a simple system of coupled LDEs for x and y

$$m\dot{y} + eBx = a \quad \text{and} \quad m\dot{x} - eBy = b$$

The system is essentially a harmonic oscillator; the solution is sinusoidal

$$x = \frac{a}{eB} + R \sin(\omega(t - t_0)) \quad \text{and} \quad y = -\frac{b}{eB} + R \cos(\omega(t - t_0))$$

where a, b, R and t_0 determined by the initial conditions. The particle makes circles of radius R in the (x, y) plane at the *cyclotron frequency*

$$\omega = \frac{eB}{m}$$

This magnetic field corresponds to the vector potential

5.1.4 Hamiltonian Mechanics and Conservation Laws

We return briefly to two of conservation laws and associated constants of motion that we introduced in Lagrangian mechanics. These conservation laws are often simple to see the Hamiltonian formalism.

- If the Hamiltonian H is not explicitly time dependent (i.e. $\frac{\partial H}{\partial t} = 0$) then H is a constant of motion (i.e. $\frac{dH}{dt} = 0$).

$$\frac{dH}{dt} = \frac{\partial H}{\partial q_i} \dot{q}_i + \frac{\partial H}{\partial p_i} \dot{p}_i + \frac{\partial H}{\partial t} = -p_i \dot{q}_i + \dot{q}_i p_i + \frac{\partial H}{\partial t} = \frac{\partial H}{\partial t}$$

The equality $\frac{dH}{dt} = \frac{\partial H}{\partial t}$ means $\frac{\partial H}{\partial t} = 0 \implies \frac{dH}{dt} = 0$.

- If an ignorable (aka cyclic) coordinate q doesn't appear in the Lagrangian, the conjugate momentum $p_q = \frac{\partial L}{\partial \dot{q}}$ is a constant of motion. Because q doesn't appear in the Lagrangian it certainly doesn't appear in the Hamiltonian, which is constructed from the Lagrangian, so $\frac{\partial H}{\partial q} = 0$. We then have

$$\dot{p}_q \equiv \frac{\partial H}{\partial q} = 0$$

so p_q is a constant of motion.

5.1.5 Principle of Least Action á la Hamilton

- In Lagrangian mechanics we defined action as

$$S = \int_{t_1}^{t_2} L(q_i, \dot{q}_i, t) dt = 0$$

and then derived the Lagrange equations of motion from the least action principle by requiring that $\delta S = 0$ for all paths with fixed endpoints so that $\delta q_i(t_1) = \delta q_i(t_2) = 0$.

- In the Hamiltonian formalism, we define action as

$$S = \int_{t_1}^{t_2} (p_i \dot{q}_i - H(q_i, p_i, t))$$

where $\dot{q}_i = \dot{q}_i(q_i, p_i)$.

We then consider varying q_i and p_i *independently*, whereas in the Lagrangian formalism a variation of q_i lead to a variation of \dot{q}_i . The independent approach makes sense: a theme of the Hamiltonian approach is treating q_i and p_i on equal footing. We have

$$\delta S = \int_{t_1}^{t_2} \left[\delta p_i \dot{q}_i + p_i \delta \dot{q}_i - \frac{\partial H}{\partial q_i} \delta q_i - \frac{\partial H}{\partial p_i} \delta p_i \right] dt$$

Adding and subtracting $\dot{p}_i \delta q_i$ in the integrand and factoring gives

$$\delta S = \int_{t_1}^{t_2} \left[\left(\dot{q}_i - \frac{\partial H}{\partial p_i} \right) \delta p_i + \left(-\dot{p}_i - \frac{\partial H}{\partial q_i} \right) \delta q_i + \{p_i \delta \dot{q}_i + \dot{p}_i \delta q_i\} \right] dt$$

- The last term in the curly brackets can be written as $\frac{d}{dt}[p_i \delta q_i]$ and integrated:

$$\delta S = \int_{t_1}^{t_2} \left[\left(\dot{q}_i - \frac{\partial H}{\partial p_i} \right) \delta p_i + \left(-\dot{p}_i - \frac{\partial H}{\partial q_i} \right) \delta q_i \right] dt + [p_i \delta q_i]_{t_1}^{t_2}$$

- Hamilton's equations await us in the large parentheses. If we look for extrema $\delta S = 0$ that hold for all δp_i and δq_i we must have

$$\left(\dot{q}_i - \frac{\partial H}{\partial p_i} \right) = 0 \quad \text{and} \quad \left(-\dot{p}_i - \frac{\partial H}{\partial q_i} \right) = 0$$

and the boundary conditions $\delta q_i(t_1) = \delta q_i(t_2) = 0$ so that the term $[p_i \delta q_i]_{t_1}^{t_2}$ vanishes and $\delta S = 0$.

5.2 Poisson Brackets

5.2.1 Definition and Properties

- Let $f(q_i, p_i)$ and $g(q_i, p_i)$ be two functions defined on phase space where $i = 1, \dots, n$. The *Poisson bracket* of the functions f and g is

$$\{f, g\} = \sum_{i=1}^n \left(\frac{\partial f}{\partial q_i} \frac{\partial g}{\partial p_i} - \frac{\partial f}{\partial p_i} \frac{\partial g}{\partial q_i} \right)$$

We generally use Einstein summation notation and drop the sum symbol \sum .

- Some properties of Poisson brackets are:
 - $\{f, g\} = -\{g, f\}$
 - Linearity: $\{\alpha f + \beta g, h\} = \alpha \{f, h\} + \beta \{g, h\}$ for all real scalars $\alpha, \beta \in \mathbb{R}$
 - Leibniz (product) rule: $\{fg, h\} = f\{g, h\} + g\{f, h\}$
 - Jacobi identity: $\{f, \{g, h\}\} + \{g, \{h, f\}\} + \{h, \{f, g\}\} = 0$

Note that the Poisson bracket obeys a similar algebraic structure as the differentiation operator d and the matrix commutator $[\cdot, \cdot]$.

- The next set of interesting identities are

$$\{q_i, q_j\} = 0 \quad \{p_i, p_j\} = 0 \quad \{q_i, p_j\} = \delta_{ij}$$

These are fairly trivial to prove, but I was confused with the proofs in my first encounter and couldn't easily find a good explanation in the literature, so here is a step-by-step walk-through for other slow learners like me. Feel free to skip. The key to all three is using the (obvious) relations

$$\frac{\partial p_i}{\partial q_j} = \frac{\partial q_i}{\partial p_j} = 0 \quad \text{and} \quad \frac{\partial p_i}{\partial p_j} = \frac{\partial q_i}{\partial q_j} = \delta_{ij}$$

and being comfortable with the implicit summation notation. I'll write the sum explicitly for the first, then follow the Einstein convention

$$\{q_i, q_j\} = \sum_{k=1}^n \left(\frac{\partial q_i}{\partial q_k} \frac{\partial q_j}{\partial p_k} - \frac{\partial q_i}{\partial p_k} \frac{\partial q_j}{\partial q_k} \right) = \sum_{k=1}^n (0 + 0)$$

Here the $\frac{\partial q}{\partial p}$ terms guarantee zeros everywhere. The second identity is completely analogous. In Einstein summation notation we have

$$\{p_i, p_j\} = \frac{\partial p_i}{\partial q_k} \frac{\partial p_j}{\partial p_k} - \frac{\partial p_i}{\partial p_k} \frac{\partial p_j}{\partial q_k} = 0$$

This time the $\frac{\partial p}{\partial q}$ terms give us zeros everywhere. Finally,

$$\{q_i, p_j\} = \frac{\partial q_i}{\partial q_k} \frac{\partial p_j}{\partial p_k} - \frac{\partial q_i}{\partial p_k} \frac{\partial p_j}{\partial q_k}$$

The $\frac{\partial q_i}{\partial p_k} \frac{\partial p_j}{\partial q_k}$ is zero everywhere. The only possible non-zero term occurs if $k = i = j$ in which case $\frac{\partial q_i}{\partial q_i} \frac{\partial p_j}{\partial p_j} = 1 \cdot 1 = 1$. If $i \neq j$ either $\frac{\partial q_i}{\partial q_i}$ or $\frac{\partial p_j}{\partial p_j}$ is zero.

- Next, a more interesting identity: for any function $f(q_i, p_i, t)$ we have

$$\frac{df}{dt} = \{f, H\} + \frac{\partial f}{\partial t}$$

We first differentiate f , then recognize the Hamilton equations $\dot{q}_i = \frac{dq}{dt} = \frac{\partial H}{\partial p_i}$ and $\dot{p} = \frac{dp}{dt} = -\frac{\partial H}{\partial q}$

$$\begin{aligned} \frac{df}{dt} &= \frac{\partial f}{\partial q_i} \frac{dq_i}{dt} + \frac{\partial f}{\partial p_i} \frac{dp_i}{dt} + \frac{\partial f}{\partial t} \frac{dt}{dt} = \frac{\partial f}{\partial q_i} \frac{\partial H}{\partial p_i} - \frac{\partial f}{\partial p_i} \frac{\partial H}{\partial q_i} + \frac{\partial f}{\partial t} \\ &= \{f, H\} + \frac{\partial f}{\partial t} \end{aligned}$$

The result is important. It tells us that any function $I = I(p_i, q_i)$ for which $\{I, H\} = 0$ is a constant of motion. Why? Simple: if $\{I, H\} = 0$ then $\frac{dI}{dt} = \frac{\partial I}{\partial t}$ and $\frac{\partial I}{\partial t} = 0 \implies \frac{dI}{dt} = 0$. If $\{I, H\} = 0$ we say that the functions I and H *Poisson commute*.

As an example, we know from Lagrangian mechanics that if q_i is a cyclic coordinate (i.e. $\frac{\partial H}{\partial q_i} = 0$) then the corresponding generalized momentum p_i is a constant motion.

Does this hold in the language of Poisson brackets? Indeed. From the identities $\frac{\partial p}{\partial q} = 0$ and $\frac{\partial p_i}{\partial p_j} = \delta_{ij}$ and then applying $\dot{p}_i \equiv 0$ we have

$$\{p_i, H\} = \frac{\partial p_i}{\partial q_j} \frac{\partial H}{\partial p_j} - \frac{\partial p_i}{\partial p_j} \frac{\partial H}{\partial q_j} = 0 - \frac{\partial H}{\partial q_i} = \dot{p}_i = 0$$

- Next, an important consequence of the Jacobi identity and $\{f, g\} = -\{g, f\}$. If I and J are both constants of motion, then

$$\{\{I, J\}, H\} = \{I, \{J, H\}\} + \{\{I, H\}, J\} = 0$$

which means $\{I, J\}$ is a constant of motion. In other words, if we know two constants of motion, we can generate more using the Poisson brackets and the Jacobi identity.

5.2.2 Poisson Brackets and Angular Momentum

- Poisson bracket show us something cool: if two components of angular momentum (e.g. L_1 and L_2) are conserved, the third component L_3 must also be conserved. This is an application of the theorem stating $\{I, J\}$ is conserved if I and J are separately conserved. We just need to show that $\{L_1, L_2\} = L_3$. We have

$$\begin{aligned} L_1 &= r_2 p_3 - r_3 p_2 & \text{and} & & L_2 &= r_3 p_1 - r_1 p_3 \\ \frac{\partial L_1}{\partial \mathbf{r}} &= (0, p_3, -p_2) & \text{and} & & \frac{\partial L_2}{\partial \mathbf{r}} &= (-p_3, 0, p_1) \\ \frac{\partial L_1}{\partial \mathbf{p}} &= (0, -r_3, r_2) & \text{and} & & \frac{\partial L_2}{\partial \mathbf{p}} &= (r_3, 0, -r_1) \end{aligned}$$

Now we just need to put the pieces together. Using the dot product gives

$$\{L_1, L_2\} = \frac{\partial L_1}{\partial \mathbf{r}} \cdot \frac{\partial L_2}{\partial \mathbf{p}} - \frac{\partial L_1}{\partial \mathbf{p}} \cdot \frac{\partial L_2}{\partial \mathbf{r}} = r_1 p_2 - r_2 p_1 = L_3$$

- Now if L_1, L_2 and L_3 are conserved, the entire vector \mathbf{L} is conserved. In other words, if two components (e.g. L_1 and L_2) are conserved, the entire vector \mathbf{L} is automatically conserved. Note also the quantum mechanic-like relationship

$$\{L^2, L_i\} = \{L_1^2 + L_2^2 + L_3^2, L_i\} = 0 \quad \text{for} \quad i = 1, 2, 3$$

- Next, we consider the Laplace-Runge-Lenz vector in the language of Poisson brackets. Recall the LRL vector is

$$\mathbf{A} = \frac{1}{m} \mathbf{p} \times \mathbf{L} - \hat{\mathbf{r}}$$

where $\hat{\mathbf{r}} = \mathbf{r}/r$. The vector satisfies $\mathbf{A} \cdot \mathbf{L} = 0$ since both $\mathbf{r} \times \mathbf{L}$ and $\hat{\mathbf{r}}$ are orthogonal to \mathbf{L} . More complicated are the Poisson bracket identities

$$\{L_a, A_b\} = \epsilon_{abc} A_c \quad \text{and} \quad \{A_a, A_b\} = -\frac{2}{m} \left(\frac{\mathbf{p}^2}{2m} - \frac{1}{r} \right) \epsilon_{abc} L_c$$

where ϵ_{abc} is the Levi-Civita symbol. For a system with a central potential $V = -\frac{1}{r}$ and thus a Hamiltonian $H = \frac{\mathbf{p}^2}{2m} - \frac{1}{r}$, the last equality is simply

$$\{A_a, A_b\} = -\frac{2H}{m} \epsilon_{abc} L_c$$

in which case $\{H, \mathbf{A}\} = 0$, meaning the LRL vector is a constant of motion in an $-\frac{1}{r}$ potential.

Kepler's elliptical orbits follow from the conservation of \mathbf{A} . The dot product $\hat{\mathbf{r}}\mathbf{A}$ gives

$$\hat{\mathbf{r}} \cdot \mathbf{A} = \frac{1}{m} \hat{\mathbf{r}} \cdot (\mathbf{p} \times \mathbf{L}) - 1 \implies \hat{\mathbf{r}} \cdot \mathbf{A} + 1 = \frac{\mathbf{L}^2}{r}$$

which is the equation for an ellipse.