

# Flares in Open Clusters with K2.

## II. Pleiades, Hyades, Praesepe, Ruprecht 147, and M67

Ekaterina Ilin<sup>1,2</sup>, Sarah J. Schmidt<sup>1</sup>, Katja Poppenhäger<sup>1,2</sup>, James R. A. Davenport<sup>3</sup>, Martti H. Kristiansen<sup>4,5</sup>, Mark Omohundro<sup>6</sup>

<sup>1</sup> Leibniz-Institute for Astrophysics Potsdam (AIP), An der Sternwarte 16, 14482 Potsdam, Germany  
e-mail: eilin@aip.de,

<sup>2</sup> Institute for Physics and Astronomy, University of Potsdam, Karl-Liebknecht-Str. 24/25, 14476 Potsdam, Germany

<sup>3</sup> Department of Astronomy, University of Washington, Seattle, WA 98195, USA

<sup>4</sup> Brorfelde Observatory, Observator Gyldenkernes Vej 7, DK-4340 Tølløse, Denmark

<sup>5</sup> DTU Space, National Space Institute, Technical University of Denmark, Elektrovej 327, DK-2800 Lyngby, Denmark

<sup>6</sup> Citizen Scientist, c/o Zooniverse, Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford, OX1 3RH, UK

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### ABSTRACT

**Context.** Magnetic fields are a key component in the main sequence evolution of low mass stars. Flares, energetic eruptions on the surfaces of stars, are an unmistakable manifestation of magnetically driven emission. The occurrence rates and energy distributions of flares trace stellar characteristics such as mass and age. But before flares can be used to constrain stellar properties, the flaring-age-mass relation requires proper calibration.

**Aims.** This work sets out to quantify flaring activity of independently age-dated main sequence stars for a broad range of spectral types using optical light curves obtained by the Kepler satellite.

**Methods.** Drawing from the complete K2 archive, we searched 3435  $\sim$  80 day long light curves of 2111 open cluster members for flares using the open-source software packages K2SC to remove instrumental and astrophysical variability from K2 light curves, and AltaiPony to search and characterize the flare candidates.

**Results.** We confirmed a total of 3844 flares on high probability open cluster members with ages from Zero Age Main Sequence (Pleiades) to 3.6 Gyr (M67). We extended the mass range probed in the first study of this series to span from Sun-like stars to mid-M dwarfs. We added the Hyades (690 Myr) to the sample as a comparison cluster to Praesepe (750 Myr), the 2.6 Gyr old Ruprecht 147, and several hundred light curves from the late K2 Campaigns in the remaining clusters. We found that the flare energy distribution was similar in the entire parameter space, following a power law relation with exponent  $\alpha \approx 2.0 - 2.1$ .

**Conclusions.** We confirmed that flaring rates declined with age, and declined faster for higher mass stars. Our results are in good agreement with most previous statistical flare studies. We conclude that some discrepancies between our results and flare studies that used rotation periods for their age estimates could be explained by sample selection bias towards more active stars, but others may hint at limitations to rotation as an age indicator.

**Key words.** Methods: data analysis, Stars: activity, Stars: flare, Stars: low-mass

## 1. Introduction

Flares are explosions that occur on the surfaces of practically all low mass stars down to the very bottom of the main sequence. We know that flares are magnetic re-connection events that lead to a change in field line topology and subsequent energy release (Priest & Forbes 2002). We can observe flares in nearly all electromagnetic bands, from radio to X-ray, and on all stars that possess a convection zone, from late F type stars to ultracool dwarfs (Schaefer et al. 2000; Benz & Güdel 2010; Gizis 2013). Stellar flares on cool stars can enhance the optical flux by multiple orders of magnitude within minutes or seconds (Haisch et al. 1991; Schmidt et al. 2019), and release energies up to  $\sim 10^{37}$  erg (Maehara et al. 2012; Davenport 2016). They typically exhibit blackbody emission at about 10 000 K (Hawley & Fisher 1992a; Kowalski et al. 2013), which is significantly hotter than the photospheres of these stars (2500 – 6500 K). The high intensity, strong contrast, and broad energy distributions

allow us to measure magnetic activity as traced by flares for a great variety of time resolved observations.

Statistical flare studies were pioneered from the ground (Lacy et al. 1976), but it was not until space missions like Kepler (Koch et al. 2010) and the Transiting Exoplanet Survey Satellite (TESS, Ricker et al. 2014) that investigating stellar ensembles that were not pre-selected for their activity was possible (Walkowicz et al. 2011). Statistical studies of stellar flaring activity can help us understand the underlying physical processes of flares (Benz & Güdel 2010), the nature and strength of stellar magnetic fields (Berger 2006; Odert et al. 2017), starspots (Davenport 2015; Howard et al. 2020), how flares relate to stellar angular momentum evolution (Mondrik et al. 2019; Howard et al. 2020), and how they affect the atmospheres of exoplanets (Lecavelier des Etangs et al. 2012; Loyd et al. 2018a; Tilley et al. 2019; Howard et al. 2019).

The present study has a focus on the relation between flares and the age of low mass stars on the main sequence. To this end,

we set out to quantify how flaring activity changes as a function of stellar mass and age.

In cool main sequence stars, the fraction of stars found flaring in the optical is known to increase down to spectral type M5 (Yang et al. 2017; Chang et al. 2020). Even beyond M5, and down to L dwarfs, flaring is ubiquitous (Stelzer et al. 2006; Robrade et al. 2010; Gizis 2013; Schmidt et al. 2015; Schmidt et al. 2016; Paudel et al. 2018). In fact, the prototype flare star, UV Ceti, is an M6 dwarf (Kirkpatrick et al. 1991). Flaring activity also decays with age, but stellar ages can only be determined indirectly. Using galactic latitude as a proxy for age, flaring activity appeared to decline with higher galactic latitude, that is, older age, for M dwarfs (Hilton et al. 2010; Walkowicz et al. 2011; Howard et al. 2019). In gyrochronology, fast rotation indicates young age (Barnes 2003), and slows down as the star ages. So the relation between rotation and flaring can be used to calibrate a flaring-age relation (Davenport et al. 2019).

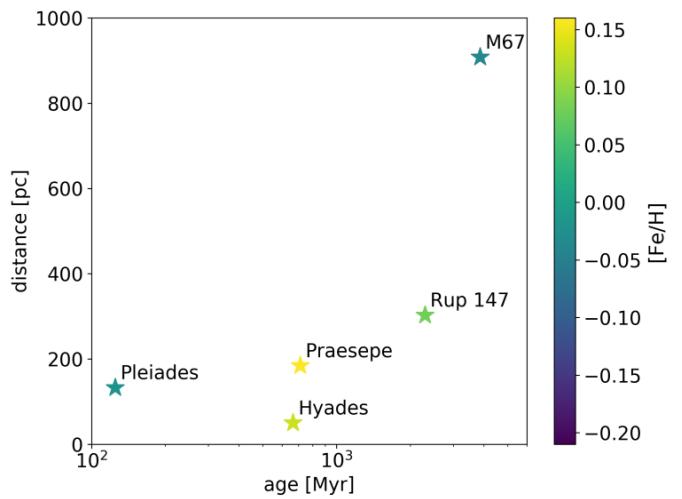
A significant portion of flaring energy is emitted in H $\alpha$  (Kowalski et al. 2013), which was also explored as an age tracer in the past (Soderblom et al. 1991; Pace 2013; Lorenzo-Oliveira et al. 2016). More recent work focused on stars older than a gigayear, and calibrated the H $\alpha$ -age relation using isochronal (Lorenzo-Oliveira et al. 2018), and asteroseismic ages (Booth et al. 2017, 2020), suggesting that an age calibration of flaring activity is possible.

While absolute stellar ages are difficult to assess, they can be differentially controlled for in coeval groups of stars. Flaring-age studies in binaries showed consistent activity for both components in the majority of targets (Lurie et al. 2015; Clarke et al. 2018). Open clusters are also coeval groups of stars with well-determined isochronal ages that have been used as a laboratory for flare studies on stars with a fixed age (Mirzoyan 1993; Chang et al. 2015). Ilin et al. (2019) (hereafter Paper I) investigated the flaring activity of late-K to mid-M dwarfs in three open clusters, the Pleiades, Praesepe, and M67, using K2 time domain photometry. We analyzed flare frequency distributions (FFDs) broken down by the stars' effective temperatures  $T_{\text{eff}}$  and ages. We found that flaring activity declined both with increasing mass and age, and that the trend was more pronounced for higher mass stars.

This study extends the results in Paper I to the age of Ruprecht 147 (2.5 Gyr), and both higher and lower masses than in the previous study. We used the now complete K2 data set, supplemented all three open clusters in Paper I with improved versions of already treated light curves in Campaigns 4 and 5, and added light curves from later Campaigns. The light curve catalog, and the determination of cluster membership and effective temperature for the investigated stars are detailed in Section 2. We describe how we used a semi-automated flare finding procedure, how we estimated flare energies, and how we parametrized the statistical properties of flares in Section 3. We present our findings in Section 4. We place our results in the context of recent flare studies, and reflect on the power law nature of FFDs in Section 5. Recently, Davenport et al. (2019) proposed an empirical parametrization of the flaring-mass-age relation based on FFDs of stars with gyrochronologically determined ages, which we also put to test in the discussion. The summary and conclusions can be found in Section 6.

## 2. Observations and data analysis

This work is based on K2 long cadence light curves with an integration time of 30 minutes that were provided by the Kepler archives hosted at the Barbara A. Mikulski Archive for



**Fig. 1.** Distance as a function of age for open clusters (star symbols) that we searched for flares, colored by [Fe/H]. Age, distance, and [Fe/H] are approximate values from a compilation of existing literature, see Appendix B.

Space Telescopes (MAST) (Section 2.1). To select a sample with independently derived ages we began by gathering open cluster membership information from the literature (Section 2.2). An overview over the cluster sample is presented in Table 1 and illustrated in Figure 1. To characterize the stars in the resulting catalog for the analysis that followed, we used multiband photometry from several all-sky surveys to assign  $T_{\text{eff}}$  (Fig. 2) and determine stellar radii (Fig. 3) using empirical relations, and eventually calculated stellar luminosities in the Kepler band (Section 2.3).

### 2.1. K2 light curves

The Kepler (Koch et al. 2010) spacecraft finished its follow-up mission K2 (Howell et al. 2014) in September 2018, after having completed nearly 20 80-day observing campaigns. With the full archive at hand, we selected and analyzed a total of 3435 light curves of high-confidence members of five open clusters. Each light curve contained up to 80 uninterrupted days of 30 min cadence observations in white light ( $4,200 - 9,000 \text{ \AA}$ ). As K2 was conducted on the two-wheeled Kepler satellite, it was subjected to substantial drift motion (spacecraft roll, Van Cleve et al. 2016) and had an overall reduced pointing accuracy. To mitigate these effects, various solutions were developed (Vanderburg & Johnson 2014; Aigrain et al. 2016; Luger et al. 2016; Luger et al. 2018). We used the K2 Systematics Correction (K2SC) pipeline (Aigrain et al. 2016) with minor modifications to de-trend the 36th data release of the K2 data products. This data release was final, and included a uniform, global reprocessing of most K2 campaigns using an improved data reduction pipeline<sup>1</sup>.

### 2.2. Open Cluster membership

For each open cluster, we obtained membership information from a selection of catalogs which we cross-matched on right ascension and declination within 3 arcsec. One

<sup>1</sup> <https://keplerscience.arc.nasa.gov/k2-uniform-global-reprocessing-underway.html>

part of the membership catalogs provided membership probabilities (Douglas et al. 2014; Cantat-Gaudin et al. 2018; Olivares et al. 2018; Reino et al. 2018; Gao 2018; Olivares et al. 2019). For the other part no probability was quoted (Rebull et al. 2016; Douglas et al. 2017; Gaia Collaboration et al. 2018a), or qualitative classifiers were given (Curtis et al. 2013; Gonzalez 2016; Rebull et al. 2017). In the latter cases we assigned approximate probabilities anchored to the set threshold for inclusion into our final sample (Appendix A). Absence in a catalog did not decrease the likelihood of membership, as each catalog had different selection biases which we did not address in this study. We set the threshold mean membership probability  $p$  for a target in our sample to 0.8. The resulting target lists were used to search the K2 archive.

All in all, we identified 2111 low mass stars in five open clusters spanning ages from ZAMS to roughly solar. Table 1 provides an overview over the final sample of stars and flares, and the adopted open cluster properties. Membership probability histograms of the final sample are displayed in Figure A.1 in Appendix A. A literature overview of age, distance, and metallicity determinations for the open clusters from which we adopt the values for this study is given in Table B.1 in Appendix B<sup>2</sup>.

### 2.2.1. Pleiades

The Pleiades, a nearby ZAMS cluster (Bell et al. 2012; Scholz et al. 2015; Dahm 2015; Yen et al. 2018; Gossage et al. 2018; Bossini et al. 2019), was observed in Campaign 4, and treated in Paper I. Here, we used 741 reprocessed K2 light curves from this open cluster, and our methodological improvements to Paper I to repeat the analysis. We revisited the members from Rebull et al. (2016), which were used in Paper I, and merged the catalog with lists of members determined by Olivares et al. (2018); Gaia Collaboration et al. (2018a); and Cantat-Gaudin et al. (2018).

### 2.2.2. Hyades

The Hyades are a 625 – 690 Myr old open cluster (Perryman et al. 1998; Salaris et al. 2004; Gossage et al. 2018; Gaia Collaboration et al. 2018a) that was observed during Campaigns 4 and 13 with K2. We merged membership tables obtained from Douglas et al. (2014); Reino et al. (2018); and Gaia Collaboration et al. (2018a) to select 170 high confidence members, of which 9 were observed in both K2 campaigns.

### 2.2.3. Praesepe

Compared across different studies, Praesepe (590 – 830 Myr, Salaris et al. 2004; Boudreault et al. 2012; Scholz et al. 2015; Yen et al. 2018; Gossage et al. 2018; Bossini et al. 2019) is a little older than the Hyades on average. The open cluster was observed during Campaign 5, and was also treated in Paper I. It was observed again during campaigns 16 and 18. We revisited the membership catalog from Douglas et al. (2014), and matched it to the members identified in Douglas et al. (2017); Rebull et al. (2017); Cantat-Gaudin et al. (2018); and Gaia Collaboration et al. (2018a). This was the richest sample with over 900 stars and nearly 2000 light curves.

<sup>2</sup> The analysis code for this section can be found at <https://github.com/ekaterinailin/flares-in-clusters-with-k2-ii> in the "Membership\_Matching" directory

### 2.2.4. Ruprecht 147

Ruprecht 147 is an open cluster with estimated ages from 2.0 Gyr to 3.1 Gyr (Curtis et al. 2013; Scholz et al. 2015; Gaia Collaboration et al. 2018a; Torres et al. 2018) observed during Campaign 7 with K2. We used the mean membership probabilities obtained from Curtis et al. (2013); Cantat-Gaudin et al. (2018); Olivares et al. (2019), and Gaia Collaboration et al. (2018a) to identify the 53 most likely members with K2 light curves.

### 2.2.5. M67

M67 is an old open cluster that has 234 members which were observed during campaigns 5, 16 and 18. Its age is about solar, with estimates ranging from 3.4 Gyr to 4.3 Gyr (Salaris et al. 2004; Önehag et al. 2011; Dias et al. 2012; Scholz et al. 2015; Barnes et al. 2016; Bossini et al. 2019). The relatively large distance of about 900 pc (Dias et al. 2012) limited the number of observable low mass stars in this otherwise rich cluster. We did not find any flares in M67 in Campaign 5 (Paper I) light curves of members identified by Gonzalez (2016). Campaigns 16 and 18 delivered both additional observations, and new targets to our previous sample, yielding almost 500 light curves for this open cluster. We merged the members from Gonzalez (2016) with a recent study based of Gaia DR2 data (Gao 2018).

## 2.3. Effective temperatures, stellar radii, and luminosities

To study the flaring-age-mass relation, we required effective temperatures ( $T_{\text{eff}}$ , Section 2.3.1), and the luminosities in the Kepler band ( $L_{Kp,*}$ , Section 2.3.3) of the investigated stars.  $L_{Kp,*}$  was required to determine the flare energies, and  $T_{\text{eff}}$  was used as a proxy to stellar mass. Since our catalog of stars spanned a wide variety of observables, no reliable Gaia parallaxes were available for a significant fraction of the targets, no single photometric survey covered all stars with broadband observations, and there was no empirical relation available that could have been used to derive  $T_{\text{eff}}$  for the entire range of spectral types. Empirical color-temperature relations (CTR) suffer from systematic errors that stem both from the different methods applied, and from sample selection biases. We therefore used several empirical relations in their appropriate ranges to obtain  $T_{\text{eff}}$  from each, and draw a more reliable mean. Targets that were lacking sufficient photometric data to derive  $T_{\text{eff}}$ , or that were too hot to be expected to have a convective envelope ( $T_{\text{eff}} \geq 7000$  K), were removed from the sample. We dropped all targets where the uncertainty on the weighted mean  $T_{\text{eff}}$  was greater than 10 %. Only targets that were assigned a  $T_{\text{eff}}$  were searched for flares. To eventually derive  $L_{Kp,*}$  from  $T_{\text{eff}}$  and model spectra we also required stellar radii as an intermediate step (Section 2.3.2).

### 2.3.1. Effective temperatures $T_{\text{eff}}$

We determined  $T_{\text{eff}}$  using broadband photometry from the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006), the Panoramic Survey Telescope and Rapid Response System Data Release 1 (PanSTARRS DR1; Chambers et al. 2016), and Gaia Data Release 2 (Gaia DR2; Gaia Collaboration et al. 2016b, 2018b) with the following quality cuts: We required the 2MASS measurements for  $J$ ,  $H$ , and  $K$  to be "A". "A" meant that measurements had  $S/N > 10$  and  $\sigma < 0.11$ . For PanSTARRS photometry, we required that the QF\_OBJ\_GOOD quality filter flag was set. SDSS and PanSTARRS  $ugrizy$  bands

**Table 1.** Open clusters.

	d [pc]	stars	LCs	flares	campaigns	age [Myr]	[Fe/H]
Pleiades	135.6 [1]	741	741	1583	4	135 $(\frac{25}{25})$ [2]	-0.04(0.03) [3]
Hyades	46.0	170	179	402	4, 13	690 $(\frac{160}{100})$ [4]	0.13(0.02) [5]
Praeseppe	185.5 [1]	913	1965	1849	5, 16, 18	750 $(\frac{3}{7})$ [6]	0.16(0.00) [5]
Ruprecht 147	305.0 [1]	53	53	9	7	2650 $(\frac{380}{380})$ [7]	0.08(0.07) [8]
M67	908.0 [9]	234	497	1	5, 16, 18	3639 $(\frac{17}{17})$ [6]	-0.10(0.08) [3]

**Notes.** The values for age, [Fe/H], and distance  $d$  are approximate values arrived at by a comparison of existing literature (see Table B.1 in the Appendix). The uncertainties are noted in parentheses. "stars" denotes the number of cluster members with membership probability  $p > 0.8$ . "LCs" and "campaigns" are the numbers of available light curves, and the K2 campaigns during which they were observed, respectively. "flares" is the number of confirmed flares found in each cluster. The quoted Hyades distance is the mean Gaia DR parallax distance.

**References.** [1] Cantat-Gaudin et al. (2018), [2] Gossage et al. (2018), [3] Conrad et al. (2014), [4] Gaia Collaboration et al. (2018a), [5] Netopil et al. (2016), [6] Bossini et al. (2019), [7] Torres et al. (2018), [8] Bragaglia et al. (2018), [9] Dias et al. (2012)

were similar but not identical. They could be converted using Table 2 in Finkbeiner et al. (2016). And for Gaia photometry, we cut at  $\text{flux}/\text{flux\_error} \geq 10$  in the  $G$ ,  $BP$ , and  $RP$  bands. Gaia astrometry solutions were only used when the Re-normalized Unit Weight Error (RUWE) was  $< 1.4$ . We also removed foreground stars using Gaia DR2 parallaxes. We corrected the 2MASS and PanSTARRS photometry from the distant clusters M67 and Ruprecht 147 for extinction using the most recent version of the dustmaps (Green 2018) package that provides 3D dust maps derived from 2MASS and PanSTARRS photometry together with Gaia distances (Green et al. 2019). When no Gaia parallax was available we used the cluster median distance instead. If an extinction value was not available for a given star we used the average extinction value of the respective cluster. We accounted for extinction in Gaia  $BP$  and  $RP$  bands using the reddening  $E(BP - RP)$  derived from Gaia photometry and parallax from Gaia DR2 (Andrae et al. 2018). We dropped targets that were too bright and would saturate the detector (Kepler magnitude  $K_p \leq 9$ ). Faint targets were not filtered out by a threshold but were removed from the sample when  $T_{\text{eff}}$  did not meet the quality criteria described above.

To determine robust  $T_{\text{eff}}$  (Fig. 2), we used CTRs from Boyajian et al. (2013) and Mann et al. (2016) (erratum to Mann et al. 2015),  $T_{\text{eff}}$  derived from Gaia DR2 using the StarHorse algorithm (Queiroz et al. 2018), and  $T_{\text{eff}}$  inferred from Gaia DR2 using the Apsis pipeline (Bailer-Jones et al. 2013; Andrae et al. 2018).

Boyajian et al. (2013) determined CTRs from a range of interferometrically characterized stars using  $g - r$ ,  $g - i$ ,  $g - z$ ,  $g - J$ ,  $g - H$ , and  $g - K$  colors from SDSS and Johnson magnitudes for A to K stars. Their sample was centered on solar metallicity, so we constrained the use of these CTRs to stars with  $-0.25 < [\text{Fe}/\text{H}] < 0.25$ . Following Boyajian et al. (2013), we transformed 2MASS  $JHK$  to  $J - H$ ,  $H - K$ , and  $J - K$  in the Johnson system from 2MASS to the Bessell-Brett system (Carpenter 2001), and from Bessell-Brett to Johnson using the relations in Bessell & Brett (1988).

Mann et al. (2015) derived CTRs from absolutely calibrated spectra to which they fitted atmospheric models to obtain  $T_{\text{eff}}$ . They provided transformations for SDSS/2MASS  $r - z$  and  $r - J$ , or Gaia  $BP - RP$  where extra information could be added from metallicity or 2MASS  $J - H$ . The relations in Mann et al. (2015) were only valid if metallicity was sufficiently close to solar, which was satisfied for the open clusters in this paper (see Table 1).

We supplemented our estimates with  $T_{\text{eff}}$  estimates from Anders

et al. (2019) who determined distances, extinctions, and various stellar parameters for 137 million stars in Gaia DR2 using the StarHorse pipeline (Queiroz et al. 2018).

Gaia DR2 published  $T_{\text{eff}}$  for over 160 million sources (Gaia Collaboration et al. 2018b). The typical uncertainty was quoted at 324 K, but it was lower for stars above  $\sim 4100$  K and below  $\sim 6700$  K, so that we adopt 175 K which is above the quoted root-median-squared error in this  $T_{\text{eff}}$  range (Andrae et al. 2018), and use provided values only in this range.

Finally, we plotted color-magnitude diagrams in all aforementioned colors, and flagged outliers from the main sequence. These stars were excluded from the analysis.<sup>3</sup>

### 2.3.2. Stellar radii

We used a catalog of empirically characterized stars (Yee et al. 2017) and empirical  $T_{\text{eff}} - R_*$  relations derived by Mann et al. (2015, 2016) to obtain  $R_*$  from the  $T_{\text{eff}}$  for the stars in our sample (Fig. 3). Yee et al. (2017) compiled 404 stars with high-resolution spectra from the literature, and performed their own observations of mid to late K-dwarfs, achieving a coverage of low mass stars from 7000 K down to 3000 K. For these stars, the resulting catalog was accurate to 100 K in  $T_{\text{eff}}$ , 15 % in  $R_*$ , and 0.09 dex in [Fe/H]. We interpolated between stars from the catalog to our derived  $T_{\text{eff}}$ , and propagated the resulting scatter to the uncertainty in  $R_*$  if  $T_{\text{eff}} > 3500$  K. For stars with  $T_{\text{eff}} < 3500$  K we used  $T_{\text{eff}} - R_*$  relations derived by Mann et al. (2015, 2016).

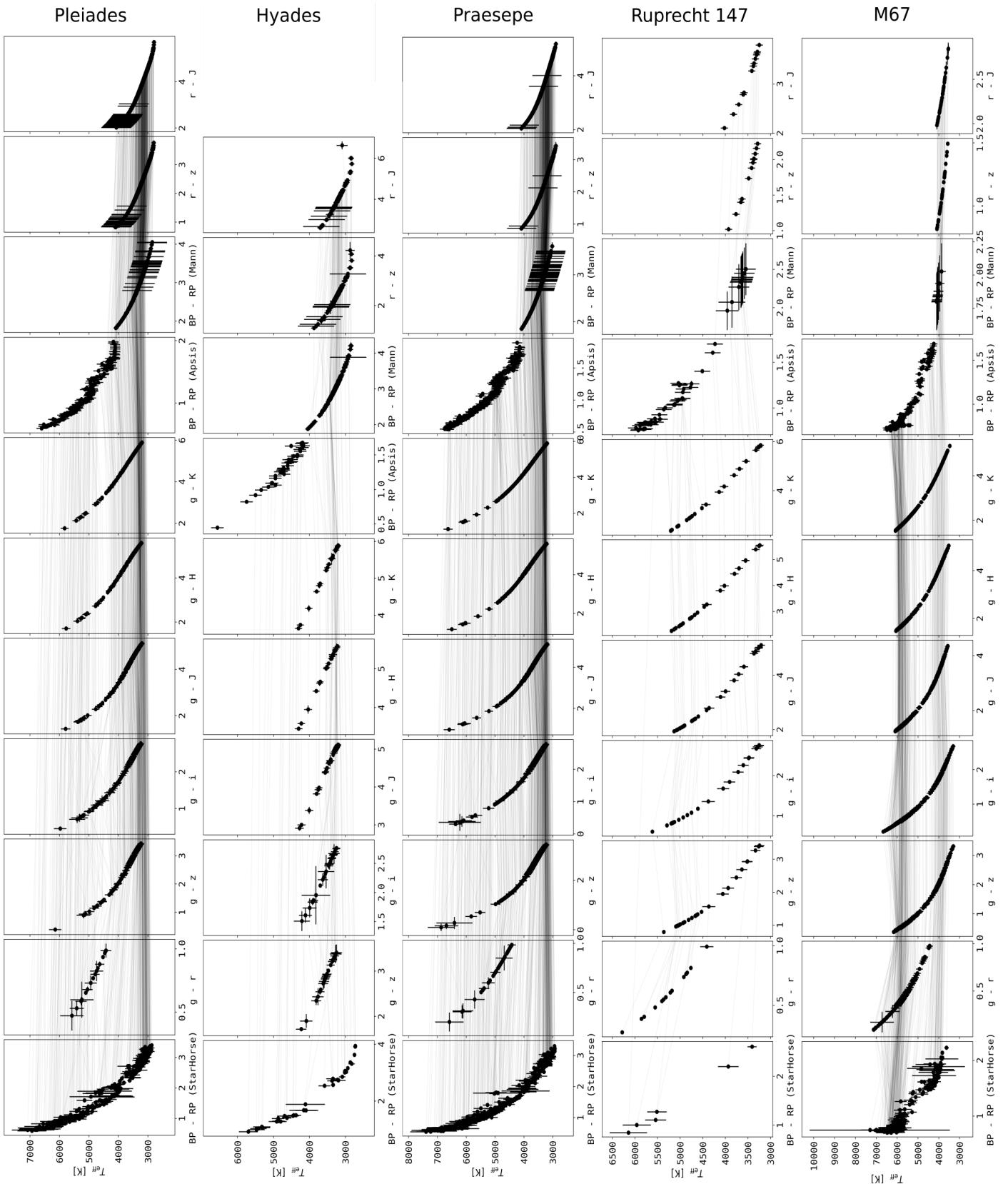
### 2.3.3. Stellar spectra and luminosities

We assigned spectra to our targets from the SpecMatchEmp<sup>4</sup> (Yee et al. 2017) and the FlareSpecPhot<sup>5</sup> libraries (Kirkpatrick et al. 2010; Burgasser et al. 2007, 2008, 2010, 2004; Cruz et al. 2004; Burgasser & McElwain 2006; Rayner et al. 2009; Doi et al. 2010; Filippazzo et al. 2015; Cruz et al. 2003; West et al. 2011; Bochanski et al. 2010, 2007; Schmidt et al. 2010, 2015, 2014; Schmidt 2014; Mann et al. 2015). When a spectrum was available for the derived spectral type in FlareSpecPhot, we preferred it over SpecMatchEmp, which was the case for all stars cooler than M0, where we

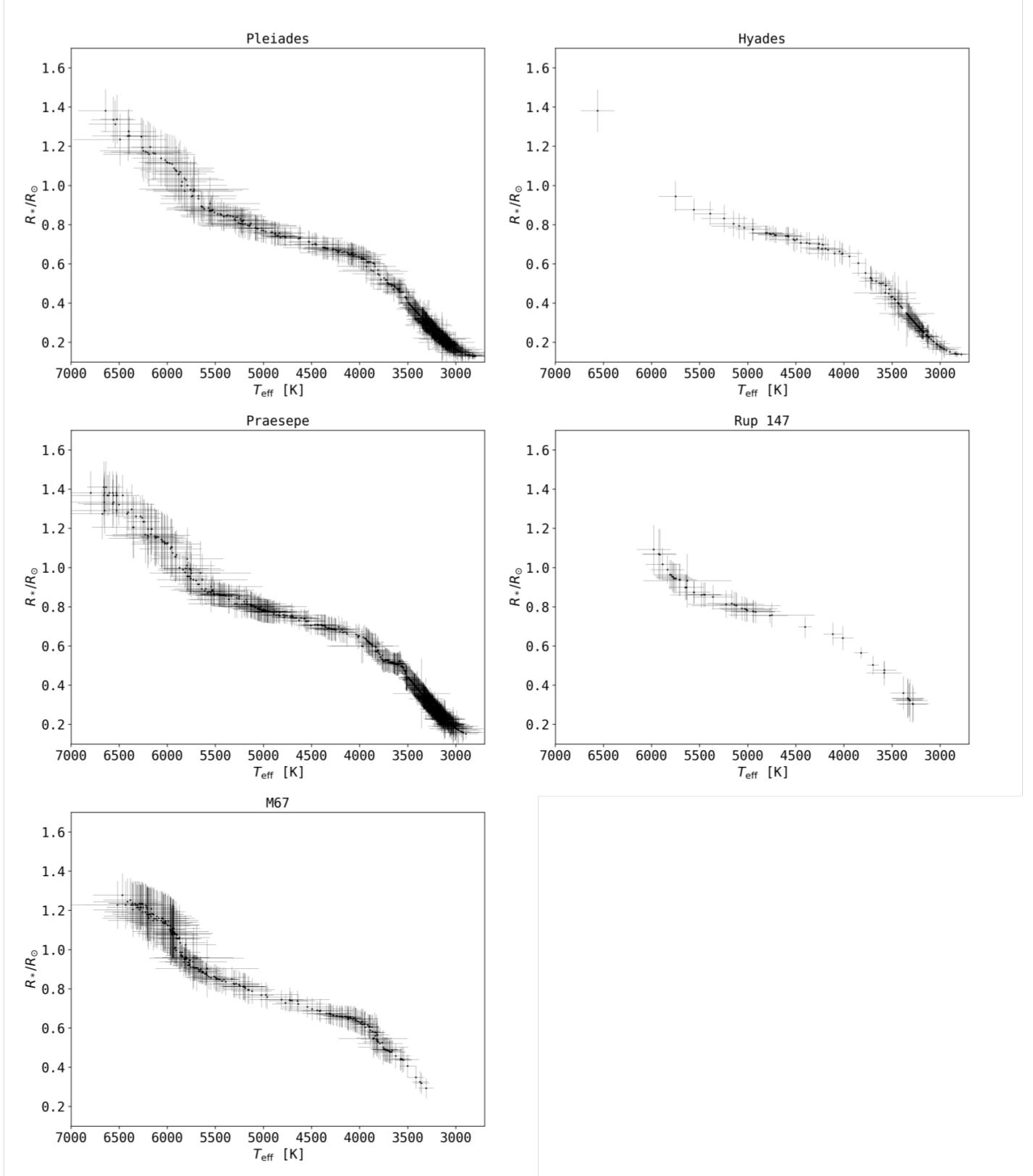
<sup>3</sup> The analysis code for this section can be found at <https://github.com/ekaterinailin/flares-in-clusters-with-k2-ii> in the "StellarParameters" directory

<sup>4</sup> <https://specmatch-emp.readthedocs.io/en/latest/>

<sup>5</sup> <https://github.com/sjschmidt/FlareSpecPhot>



**Fig. 2.** Empirical color- $T_{\text{eff}}$  relations applied to high probability members of the five open clusters in this study. Individual stars are connected with lines between sub-plots to illustrate differences in the relations. We used PanSTARRS  $griz$ , 2MASS  $JHK$ , and Gaia DR2  $BP$  and  $RP$  photometry with the following relations:  $BP - RP$  (StarHorse, Anders et al. 2019);  $g - r$ ,  $g - i$ ,  $g - z$ ,  $g - J$ ,  $g - H$ ,  $g - K$  (Boyajian et al. 2013);  $BP - RP$  (Apsis, Andrae et al. 2018);  $BP - RP$  (Mann),  $r - z$ , and  $r - J$  (Mann et al. 2016). The majority of connecting lines show very little slope, indicating that different CTRs are consistent with one another when applied to our sample. As expected, some relations show noticeable systematic differences, especially at the high and low temperature ends of their scopes of validity.



**Fig. 3.**  $T_{\text{eff}} - R_*$  distributions for all clusters in this study. Each panel shows the distribution of high probability members of the open cluster indicated in the panel title. Stellar radii  $R_*$  are given in units of solar radii. Uncertainties on  $T_{\text{eff}}$  and  $R_*$  (grey error bars) were propagated from the photometry, the color- $T_{\text{eff}}$  relations, and empirical  $T_{\text{eff}} - R_*$  relations. The uncertainties on  $R_*$  were typically  $\sim 10 - 15\%$ , which matches the scatter in the empirical  $T_{\text{eff}} - R_*$  relations. The uncertainties on  $T_{\text{eff}}$  roughly correspond to the size of the temperature bins that we chose in this work (see, for example, Fig. 5).

mapped spectral type to effective temperature (Pecaut & Mamajek 2013). We then combined  $R_*$ ,  $T_{\text{eff}}$ , and spectra to projected bolometric luminosities  $L_{\text{bol},*}$ , and projected luminosities in the Kepler band  $L_{\text{Kp},*}$  (Shibayama et al. 2013; Paper I. Uncertainties on  $L_{\text{Kp},*}$  ranged from 9 % to 52 % with a median value of 17 %.

### 3. Methods

We developed the open source software AltaiPony<sup>6</sup> as a toolbox for statistical flare studies tailored but not limited to time series observations obtained by Kepler, K2, and TESS. Its functionality includes light curve de-trending, flare finding and characterization, injection- and recovery of synthetic flares, visualization commands, and methods for the statistical analysis of flare samples. We used AltaiPony to detect flare candidates (Section 3.1), and determine their equivalent durations (Section 3.2). We validated all events by eye, and calculated their energies in the Kepler band (Section 3.3). The flare rates are believed to follow a power law distribution that spans multiple orders of magnitude in energy. We used AltaiPony’s `FFD.fit_mcmc_powerlaw()` method to fit the power law parameters simultaneously using the Markov Chain Monte Carlo method to sample from the posterior distribution (Section 3.4).

#### 3.1. Flare finding

We used the open source software AltaiPony to automatically detect and characterize flares in our sample. The code base relied on K2SC<sup>7</sup> (Aigrain et al. 2016) to remove instrumental and astrophysical variability from K2 light curves. We did not use the de-trended light curves available in MAST<sup>8</sup>. Instead, we used K2SC to de-trend light curves from the high quality, uniformly processed and documented final K2 data release ourselves. We clipped outliers at  $3\sigma$  iteratively, as compared to the original work, where outliers were clipped at  $5\sigma$  (Aigrain et al. 2016). We also masked  $\sqrt{n}$  data points before and after the sequences of outliers with lengths  $n > 1$ , because these regions often showed physically elevated flux levels below the  $3\sigma$  threshold.

After de-trending, the flare finder algorithm searched for continuous observing periods, defined as time series that were longer than 10 data points at a minimum cadence of 2 h. All further routines were run on these observing periods. We estimated the scatter in the de-trended flux using a rolling standard deviation with a 7.5 h window excluding the outliers. In the next step, the finder iteratively clipped excursions from the median value at  $3\sigma$ . After each iteration, outliers were replaced with the last median value. Either after convergence, or 50 iterations, the resulting median value was adopted.

The previously masked outliers in the light curve were assigned the mean uncertainty from the remainder of the light curve. Thus we used the non-flaring parts of the light curve to define the scatter in the flaring parts. This procedure assumed that the scatter of the 80 day light curve was stable throughout the observation. In K2, stability could not be guaranteed, leading to false positive detections in regions of light curves where the uncertainty was underestimated due to strong systematic effects that K2SC could not compensate. These false positives were later removed manually.

Using the iteratively derived median as quiescent flux and the estimated scatter, we flagged all outliers above  $3\sigma$  in the residual light curve. For flare candidates, we required at least three consecutive data points to fulfil the  $3\sigma$ -criterion. We merged the candidates into single candidate events if they were less than four data points apart.

With the automated procedure we detected 6916 flare candidates. However, the Kepler flare sample has shown to be difficult to treat in a fully automated way. Without manual vetting, the event samples remained significantly contaminated (Yang & Liu 2019). As K2 was subject to severe technical difficulties, we expected that the contamination would be even higher. Some light curves could not be de-trended using K2SC alone. Light curves with extreme astrophysical signal like deep transits, rotational modulation on time scales of a few hours or passages of bright solar system objects (SSO) had to be masked or fitted with an additional sinusoidal component to the K2SC-treated time series. Other SSOs mimicked the fast rise and exponential decay shape of flares, and could only be identified using an extension to the Virtual Observatory (VO) tool sky body tracker (SkyBoT, Berthier et al. 2006; Berthier et al. 2016). We identified 144 SSOs using SkyBot and visual inspection with `lightkurve` (Lightkurve Collaboration et al. 2018), that is 3.6 % of all confirmed flare candidates (see Appendix C). A number of light curves were excluded from the flare search because they saturated the detector, or because the target aperture overlapped with broken pixels. Some very faint targets and extreme variables could not be searched because the de-trending did not terminate successfully. The online version of the final flare sample (Table 2) includes explanatory flags, and notes on the excluded targets. Eventually, we vetted all candidates manually and confirmed 3844 events.

#### 3.2. Equivalent duration ED

For each candidate, flaring time, amplitude and equivalent duration ( $ED$ ) were returned.  $ED$  is the area between the LC and the quiescent flux, that is, the integrated flare flux divided by the median quiescent flux  $F_0$  of the star, integrated over the flare duration (Gershberg 1972):

$$ED = \int dt \frac{F_{\text{flare}}(t)}{F_0}. \quad (1)$$

$ED$  is a quantity independent of stellar calibration and distance that is suited to compare flaring activity on stars where these properties are uncertain.  $ED$  describes the time during which the non-flaring star releases as much energy as the flare itself. This time can be shorter or longer than the actual flare duration. The uncertainty in  $ED$  depends on the light curve noise properties, time sampling, spacecraft roll, and potentially other systematic effects introduced by the de-trending procedure.

#### 3.3. Kepler flare energy $E_{\text{Kp}}$

Multiband time-resolved observations of active M dwarfs have shown that thermal continuum flux accounts for the majority of the energy radiated by flares at optical wavelengths (Kowalski et al. 2013). The effective temperature of this blackbody varies: While solar flares are relatively cool, with  $T_{\text{eff}} \approx 5\,000 - 7\,000$  K (Kleint et al. 2016; Kerr & Fletcher 2014; Watanabe et al. 2013; Namekata et al. 2017), most stellar events emit in the blue, and exhibit temperatures of about  $9\,000 - 10\,000$  K (Hawley & Fisher 1992b; Kretzschmar 2011;

<sup>6</sup> [altaipony.readthedocs.io](https://altaipony.readthedocs.io)

<sup>7</sup> [github.com/OxEStech/k2sc](https://github.com/OxEStech/k2sc)

<sup>8</sup> [archive.stsci.edu/prepds/k2sc/](https://archive.stsci.edu/prepds/k2sc/)

Davenport et al. 2012; Shibayama et al. 2013). However, in ultraviolet time resolved spectral observations, one superflare was reported at a blackbody temperature of 15 500 K (Loyd et al. 2018b), and another could even be fit with a single 40 000 K spectral energy distribution (Froning et al. 2019). A dependence of flare temperature on stellar age, or mass, or both, would additionally bias our analysis. For example, at about 6 200 K, the Kepler pass band captures the largest flux fraction, at 10 000 K 72 %, at 40 000 K only 4% of this value is transmitted. Given the uncertainties on stellar luminosity in the Kepler band (see Sec. 2.3.3) and on equivalent duration, the flare energies

$$E_{\text{Kp}} = L_{\text{Kp},*} \cdot ED \quad (2)$$

will substantially deviate from the true released energy. If these uncertainties did not affect all flares in a similar fashion, the present analysis will have suffered from non-uniform biases that affected the flare frequency distributions on stars of different ages and temperatures, and skewed the flare distributions within each subsample. However, the comparison to other studies (see Sec. 5.1) suggested that our results were mostly consistent with studies based on similar data (see Fig. 11, and Lin et al. 2019 therein for an example), although they used different methods to infer  $E_{\text{flare}}$ .

### 3.4. Power law fits

Flare frequency distributions (FFDs) follow power law relations that cover several orders of magnitude, from solar microflares to stellar superflares (see Fig. 9 in Shibayama et al. 2013). In the cumulative distribution the frequency  $f(> E)$  of flares above a certain energy  $E$  is defined as

$$f(> E) = \frac{\beta}{\alpha - 1} E^{-\alpha+1}, \quad (3)$$

and, analogously, for  $ED$  (Gershberg 1972). We used and compared two approaches to fitting  $\alpha$  and  $\beta$  to the data. The first was a Modified Maximum Likelihood Estimator (MMLE) for power law distributions (Maschberger & Kroupa 2009). The second approach was more specifically tailored to flaring activity. Following Wheatland (2004), we used the MCMC method to sample from the predictive distribution of flares from a prototype flare source. This source has two characteristics. Firstly, it produces flare events as a homogeneous Poisson process in time. Secondly, the flare energies are power law distributed. While the MMLE was computationally efficient, and useful to obtain first results (see Appendix D), only the predictive distribution allowed us to fit for  $\alpha$  and  $\beta$  simultaneously, and determine their uncertainties, so we adopted the latter.

The posterior distribution in Wheatland (2004) captured both the Poissonian distribution of flare events in time, and their power law distribution in energy, simultaneously. The authors derived this model to be able to predict the flaring rate above a given energy for any active region on the Sun including changes in flaring activity rates as the active region evolves, and also characteristics of the active region itself, such as sunspot classifiers. In our simplification of the model, we assumed that the flare generating process did not change within the observation time span in any star in our sample ( $M = M'$  in Eq. 24 in Wheatland 2004). Another assumption was that this process was the same for all stars in the sample ( $\Lambda_{MC} = 1$  in aforementioned Eq. 24). Under these assumptions the samples of flares found in the light curves of different stars and light curves obtained during different campaigns could be stacked

together. With these simplifications to Eq. 24, we defined the joint posterior distribution for the probability  $\epsilon$  that a flare with  $ED$  or  $E_{\text{Kp}}$  above some value  $S_2$  would occur within a time period  $\Delta T$ :

$$\begin{aligned} p(\epsilon, \alpha) = & C \cdot (-\ln(1 - \epsilon))^M \\ & \cdot (\alpha - 1)^M \cdot \Gamma(\alpha) \left[ \frac{(S_2/S_1)^{M+1}}{\pi} \right]^\alpha \\ & \cdot (1 - \epsilon)^{(T/\Delta T) \cdot (S_2/S_1)^{\alpha-1} - 1}. \end{aligned} \quad (4)$$

$C$  was the normalisation constant,  $M$  was the number of events,  $T$  the total observation time.  $\Gamma$  contained the prior distribution for  $\alpha$ , and  $S_1$  denoted the detection threshold above which all flares were detected.  $\pi$  encapsulated the flare energies as

$$\pi = \prod_{i=1}^M \frac{s_i}{S_1}, \quad (5)$$

where  $\{s_1, s_2, \dots, s_m\}$  were the flare energies  $E_{\text{Kp}}$  or  $ED$ . We chose  $\Delta T$  to be equal to the total observation time of the FFD, and  $S_2$  was set to be ten times the maximum observed energy in that FFD.

From the posterior distribution of  $\epsilon$  we derived  $\beta$  by combining Poisson statistics

$$\epsilon = 1 - e^{(-f \cdot \Delta T)} \quad (6)$$

and the cumulative power law function in Eq. 3:

$$\beta = -\frac{\ln(1 - \epsilon) \cdot (\alpha - 1)}{\Delta T} \cdot S_2^{\alpha-1} \quad (7)$$

With a uniform prior for  $\alpha$  the results from the MMLE and Markov Chain Monte Carlo (MCMC) sampling from the posterior distribution were consistent within uncertainties (see Appendix D). However, the MCMC method allowed us to fit for both  $\epsilon$  and  $\alpha$  simultaneously, and to use more informative priors. The power law exponent determined for flare frequency distributions has consistently been found to be independent of age (Davenport et al. 2019), and spectral type for solar-type and low mass dwarfs (e.g., Chang et al. 2015; Howard et al. 2019; Lin et al. 2019; see Fig. 13 for an overview). We chose our prior to reflect this result: Starting from a uniform prior for  $\alpha$  and  $\epsilon$ , we found a Gaussian distribution to be an excellent fit to the posterior distribution for  $\alpha$  for the full sample of flares in  $E_{\text{Kp}}$  and  $ED$  space. We then used this Gaussian distribution as a prior for the FFDs in individual age and  $T_{\text{eff}}$  bins.

### 3.5. Injection and recovery of synthetic flares

Energies of flares found in Kepler 30 minute cadence light curves are underestimated by about 25 % compared to 1 minute cadence (Yang et al. 2018). Moreover, the low energy end of FFDs is incomplete because not all small flares of a given energy are recovered (Davenport 2016). In Paper I we injected and recovered synthetic flare signal based on the empirical flare template from Davenport et al. (2014) to determine the energy bias, and the recovery probability as a function of energy. We injected flares into the K2 light curves after de-trending them with K2SC to test the performance of the flare finding and characterization algorithms. Here, we performed this procedure on a random selection of light curves from our sample before applying K2SC. The injected flare candidates thus had to survive the K2SC de-trending in addition to flare

search prescription. As this procedure was computationally very expensive, we did not repeat the process for all light curves in our sample. We noted that both energy bias and recovery probability were better described as a function of amplitude and duration than of energy alone, and that they varied strongly from light curve to light curve. Therefore, energy bias and recovery probability should be quantified light curve by light curve. While this was not possible here, we recommend using the `FlareLightCurve.sample_flare_recovery()` method and associated visualization functions in `AltaiPony` with light curves from the original Kepler mission, and TESS, to assess the effects of light curve quality, and de-trending and flare search methods on the detected ensembles of flares.<sup>9</sup>. Aschwanden (2015) suggested that the observed incompleteness may be partly cancelled by background contamination from, for instance, cosmic rays, or solar system objects (SSO). While the effect of cosmic rays is currently not well constrained (Paper I), we could quantify the contamination by catalogued SSOs to 3.6 % (see Section 3.1).

## 4. Results

The core objective of this work was to quantify how the previously noted decline in flaring activity with age would unfold for different spectral types. To this end, we searched the long cadence light curves of stars across a broad range of spectral types in five different open clusters for flares, and measured their energies and occurrence rates. We found 3844 flares on 2111 stars in five open clusters for  $T_{\text{eff}}$  between 2500 K and 6000 K. The flares' amplitudes  $a$ , start and end cadence numbers  $c_0$  and  $c_1$ ,  $ED$ , energies  $E_{\text{Kp}}$ , along with the stars' EPIC ID,  $T_{\text{eff}}$ , and projected bolometric and Kepler band luminosities  $L_{\text{bol},*}$  and  $L_{\text{Kp},*}$  are listed in Table 2. We fitted power law relations to the flare frequency distributions (FFD) of stars binned by age, and  $T_{\text{eff}}$  (Figs. 5, 6 and Table 4). We found that flaring activity decreased with increasing age from ZAMS to 3.6 Gyr, and from mid-M to G dwarfs. Except for the stars in the coolest temperature bin (M5.5–M8, 2500–3000 K), stellar flaring activity at a given age was higher for cooler stars.

### 4.1. Flare frequency distributions

We fitted power law parameters  $\alpha$  and  $\beta$  (Eq. 3) to the full sample FFDs in  $E_{\text{Kp}}$  and  $ED$  space by sampling from the joint posterior distribution (Eq. 4), where we chose uninformative priors for both  $\alpha$  and  $\epsilon$ . Our FFDs combine the flare rates found in multiple stars with different observation times and detection thresholds that cause deviations from a powerlaw distribution at the high and low energy ends of the FFDs, respectively.

The detection threshold in any FFD was not a unique function of  $E_{\text{Kp}}$  or  $ED$ , but depended on the flares' amplitudes and durations. Since it was computationally not feasible to perform the injection and recovery of synthetic flares to determine the effect of K2SC de-trending, we could not correct for the incompleteness of flare detection at the low energy end (see Section 3.5). For the power law fit, our approximation to the detection threshold was to use only the portion of the FFD above the highest minimum energy out of all FFDs of individual stars. In the full sample (Fig. 4), the number of flares was large enough to allow us to double this energy threshold. This mitigated but did not solve the problem of a detection probability  $< 1$  that still affected the shape of the FFDs at the lower energy end.

We did not exclude even more flares on the low-energy end because it would have given increasing weight to the high-energy tail which steepened due to the effects of limited observation times of the flaring stars. The co-adding of flare samples from stars of different activity levels aggravated the high-energy incompleteness. Consequently the high-energy tail of the FFD was also incomplete. The absolute value of  $\alpha$  obtained in this study should be interpreted in light of these limitations. It remained useful as a relative anchor for the comparison within the sample assuming that similar caveats applied to all our age- and  $T_{\text{eff}}$ -binned FFDs.

We fitted a Gaussian distribution to the posterior distribution for  $\alpha$  from the full sample. The values for the full sample FFDs in  $E_{\text{Kp}}$  and  $ED$  space (Fig. 4) are summarized in Table 3. In the table,  $\alpha$  and  $\beta$  are the power law parameters from Eq. 3. The number of flares  $n_{\text{tot}}$  in the distribution is higher than the number  $n_{\text{fit}}$  of flares used to fit the power law because we omitted the low-energy tail of the distribution.

The aim of this work was to investigate the effects of age on the flaring activity for low mass stars. We therefore split up the full sample into  $T_{\text{eff}}$  bins and constructed the FFDs cluster by cluster (Figs. 5 and 6). The full sample posteriors for  $\alpha$  were then used as priors for the power law fits to the FFDs in the individual bins. The power law fit parameters to these FFDs, broken down by cluster and  $T_{\text{eff}}$ , are summarized in Table 4. The indices  $s$  and  $erg$  to  $\alpha$  and  $\beta$  indicate the fits to the FFDs in  $ED$  and  $E_{\text{Kp}}$ , respectively.  $n$  is the number of flares detected in each  $T_{\text{eff}}$  bin. We provide Table 4 to enable future science to build a comprehensive flare rate evolution model.

### 4.2. Power law intercept $\beta_s$

Relative flaring activity levels are best described by  $\beta_s$ , the power law fit intercept for the FFDs in  $ED$  space ( $R_{1s}$  in Davenport et al. 2019). Because of the definition of  $ED$  as relative to the quiescent flux of the star (Eq. 1)  $\beta_s$  is a measure of flaring activity in which the different energy budgets of stars are controlled for. We could therefore use  $\beta_s$  to compare across both  $T_{\text{eff}}$  and age at the same time. Fig. 7 can also be interpreted as a synthesis of the FFDs presented in Fig. 5. Overall,  $\beta_s$  declined with age in all investigated  $T_{\text{eff}}$  bins but the 2500 – 3000 K bin, where it increased from ZAMS to Hyades age. This reverse was also present in a K2 short cadence K7–M6 dwarf flaring study (Raetz et al. 2020). In the lowest mass bin (M5–M6) the authors measured higher flaring activity for relatively slow rotators ( $P_{\text{rot}} > 2$  d) than for their fast counterparts ( $P_{\text{rot}} < 2$  d). Mondrik et al. (2019) found evidence of increased flaring activity at intermediate rotation rates between fast and slow sequence (Barnes 2003). This finding could not be reproduced in the data obtained from the all-sky EvryFlare survey (Howard et al. 2020). We found  $\beta_s$  to decrease with age from ZAMS to 700 Myr for all stars except those with spectral types later than M5 (Fig. 7 a.). But this does not exclude that  $\beta_s$  in fact increased between these two ages before decreasing relatively rapidly within a short period of time.

### 4.3. Flaring luminosity FA

We can relate the luminosity in flares in the Kepler band to the quiescent bolometric luminosity of the star, and by this define the fractional flare luminosity  $FA$  (similar to Paper I):

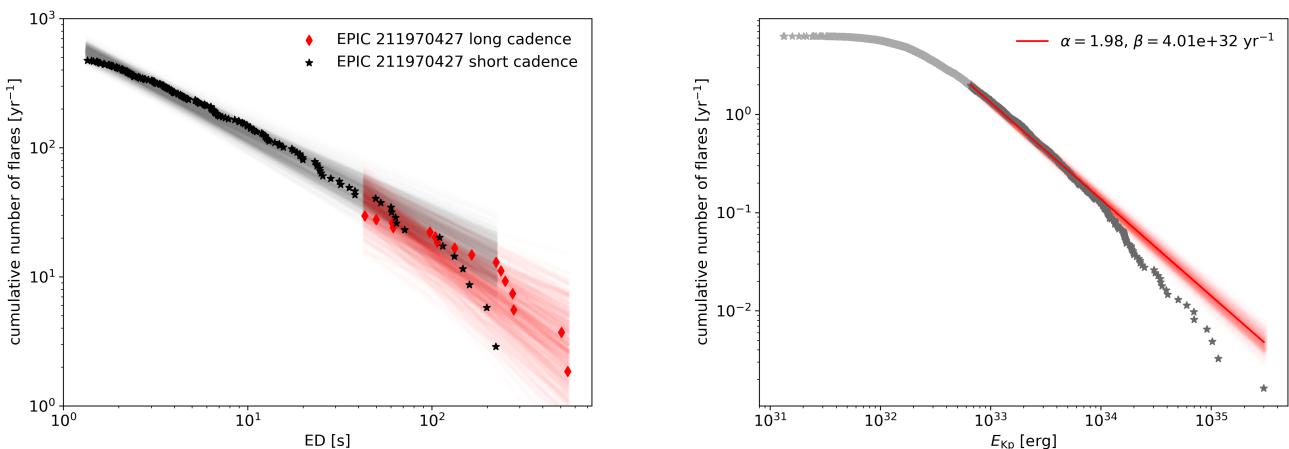
$$FA = \frac{E_{\text{Kp,flare,tot}}}{t \cdot L_{\text{bol},*}}, \quad (8)$$

<sup>9</sup> Examples of usage can be found in the software documentation.

**Table 2.** Confirmed flare candidates detected in open cluster stars observed by K2, sorted by amplitude  $a$ .

EPIC	C	cluster	$c_0$	$c_1$	$a$	$T_{\text{eff}}$ [K]	$ED$ [s]	$L_{\text{bol},*}$ [erg/s]	$L_{\text{Kp}}$ [erg/s]
211079830	4	Pleiades	105984	105993	15.330194	3097(87)	63340(63)	8.29(3.34)10 <sup>29</sup>	3.32(1.34)10 <sup>29</sup>
210720772	4	Pleiades	107181	107184	9.674757	3104(86)	19901(7)	8.66(3.45)10 <sup>29</sup>	3.48(1.38)10 <sup>29</sup>
247523445	13	Hyades	143106	143109	8.260956	2964(49)	16615(11)	4.33(1.09)10 <sup>29</sup>	1.69(0.43)10 <sup>29</sup>
212021131	5	Praesep	108974	108980	7.421916	3215(68)	19828(175)	2.01(0.65)10 <sup>30</sup>	8.22(2.65)10 <sup>29</sup>
210978953	4	Pleiades	106762	106770	6.769888	3050(95)	39467(125)	6.20(2.63)10 <sup>29</sup>	2.46(1.04)10 <sup>29</sup>
211913613	16	Praesep	156845	156849	6.690356	3218(66)	21632(147)	2.05(0.65)10 <sup>30</sup>	8.37(2.63)10 <sup>29</sup>
211127297	4	Pleiades	106754	106759	6.449569	3147(86)	20830(5)	1.13(0.45)10 <sup>30</sup>	4.59(1.81)10 <sup>29</sup>
211681193	5	Praesep	108116	108120	5.570804	3182(76)	11198(95)	1.64(0.58)10 <sup>30</sup>	6.68(2.37)10 <sup>29</sup>
211024798	4	Pleiades	104822	104826	5.394897	3290(62)	12524(63)	2.58(0.75)10 <sup>30</sup>	1.06(0.31)10 <sup>30</sup>
211134185	4	Pleiades	103891	103896	4.951629	3127(96)	17459(33)	9.99(4.33)10 <sup>29</sup>	4.03(1.75)10 <sup>29</sup>
211095280	4	Pleiades	106283	106287	4.796979	3138(93)	12186(10)	1.06(0.45)10 <sup>30</sup>	4.32(1.82)10 <sup>29</sup>
211022535	4	Pleiades	104262	104267	3.994973	2953(76)	15164(57)	3.50(1.10)10 <sup>29</sup>	1.37(0.43)10 <sup>29</sup>
211010517	4	Pleiades	106680	106685	3.968907	3252(70)	16241(171)	2.09(0.68)10 <sup>30</sup>	8.59(2.77)10 <sup>29</sup>
210846442	4	Pleiades	104410	104415	3.959977	3311(79)	13515(10)	2.88(0.96)10 <sup>30</sup>	1.18(0.40)10 <sup>30</sup>
212017838	5	Praesep	111183	111192	3.671037	3307(89)	10274(10)	3.35(1.22)10 <sup>30</sup>	1.38(0.50)10 <sup>30</sup>
211984058	5	Praesep	109952	109965	3.320636	3124(97)	16063(149)	1.17(0.51)10 <sup>30</sup>	4.73(2.06)10 <sup>29</sup>
211912899	5	Praesep	110700	110706	3.257997	3133(81)	7967(16)	1.24(0.47)10 <sup>30</sup>	5.01(1.89)10 <sup>29</sup>
211151674	4	Pleiades	106457	106467	3.166323	3072(93)	17843(106)	7.10(2.99)10 <sup>29</sup>	2.84(1.19)10 <sup>29</sup>
211822895	5	Praesep	107809	107812	3.164494	3005(177)	6883(31)	5.68(3.98)10 <sup>29</sup>	2.25(1.57)10 <sup>29</sup>
211760567	5	Praesep	109962	109965	3.061417	3245(58)	7838(9)	2.39(0.69)10 <sup>30</sup>	9.84(2.83)10 <sup>29</sup>
211939350	16	Praesep	155018	155022	2.892939	3149(150)	9542(68)	1.36(0.86)10 <sup>30</sup>	5.52(3.49)10 <sup>29</sup>
211137806	4	Hyades	106768	106775	2.482828	3127(58)	9208(8)	1.16(0.35)10 <sup>30</sup>	4.69(1.40)10 <sup>29</sup>
210674207	13	Hyades	141708	141711	2.419586	3210(74)	5343(3)	1.90(0.65)10 <sup>30</sup>	7.78(2.66)10 <sup>29</sup>
211994910	5	Praesep	109858	109864	2.401426	3325(95)	7807(14)	3.71(1.40)10 <sup>30</sup>	1.53(0.58)10 <sup>30</sup>
211010517	4	Pleiades	106561	106564	2.341076	3252(70)	4838(36)	2.09(0.68)10 <sup>30</sup>	8.59(2.77)10 <sup>29</sup>
211984058	5	Praesep	108584	108587	2.155468	3124(97)	4817(13)	1.17(0.51)10 <sup>30</sup>	4.73(2.06)10 <sup>29</sup>
211983544	18	Praesep	162792	162798	2.096077	3159(98)	6924(55)	1.44(0.63)10 <sup>30</sup>	5.88(2.55)10 <sup>29</sup>
210988354	4	Pleiades	106012	106017	2.057812	3661(144)	5725(1)	1.40(0.31)10 <sup>31</sup>	6.06(1.32)10 <sup>30</sup>
210886447	4	Pleiades	103871	103877	2.057490	3048(105)	8710(49)	6.13(2.83)10 <sup>29</sup>	2.43(1.12)10 <sup>29</sup>
211098921	4	Pleiades	105254	105262	2.041041	3300(101)	7499(5)	2.72(1.10)10 <sup>30</sup>	1.12(0.45)10 <sup>30</sup>

**Notes.** The full table is available via CDS. See Section 4 for explanations on the columns. Uncertainties are given in parentheses.

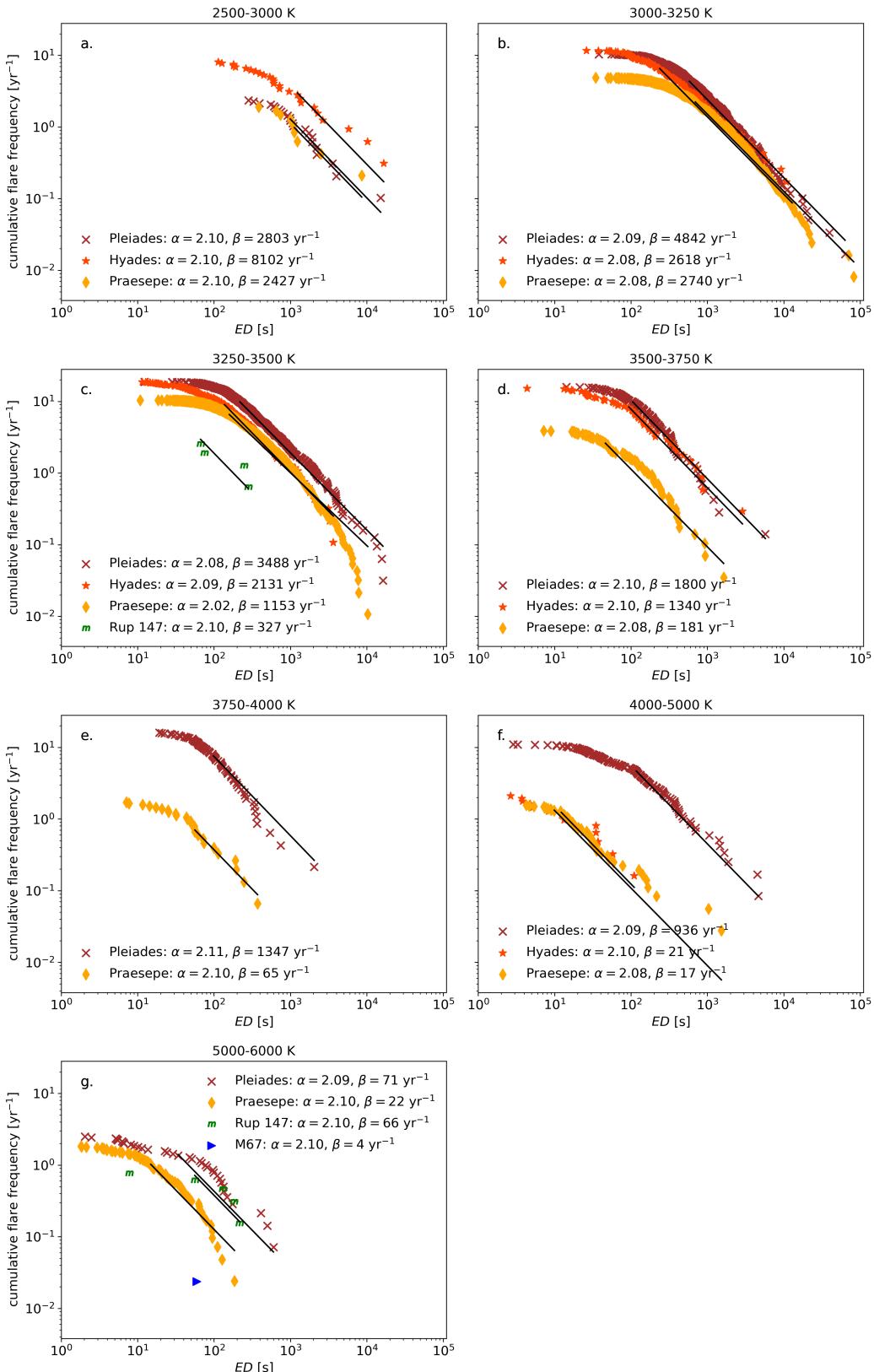


**Fig. 4.** Cumulative flare frequency distributions (FFD, grey stars) in  $ED$  (left panel) and  $E_{\text{Kp}}$  (right panel) with respective power law fits (red lines) to the full sample of confirmed flare candidates. The complete distribution (light grey stars) is larger than the portion used for the fit (dark grey stars) because we omitted the incomplete low energy tail. The red shade consists of a randomly drawn sample from the posterior distribution.  $\alpha$  and  $\beta$  are the best fit power law parameters as defined in Eq. 3. The posterior distributions were fitted with Gaussian functions that we then used as priors for the FFD fits in the  $T_{\text{eff}}$  bins in Figs. 5 and 6.

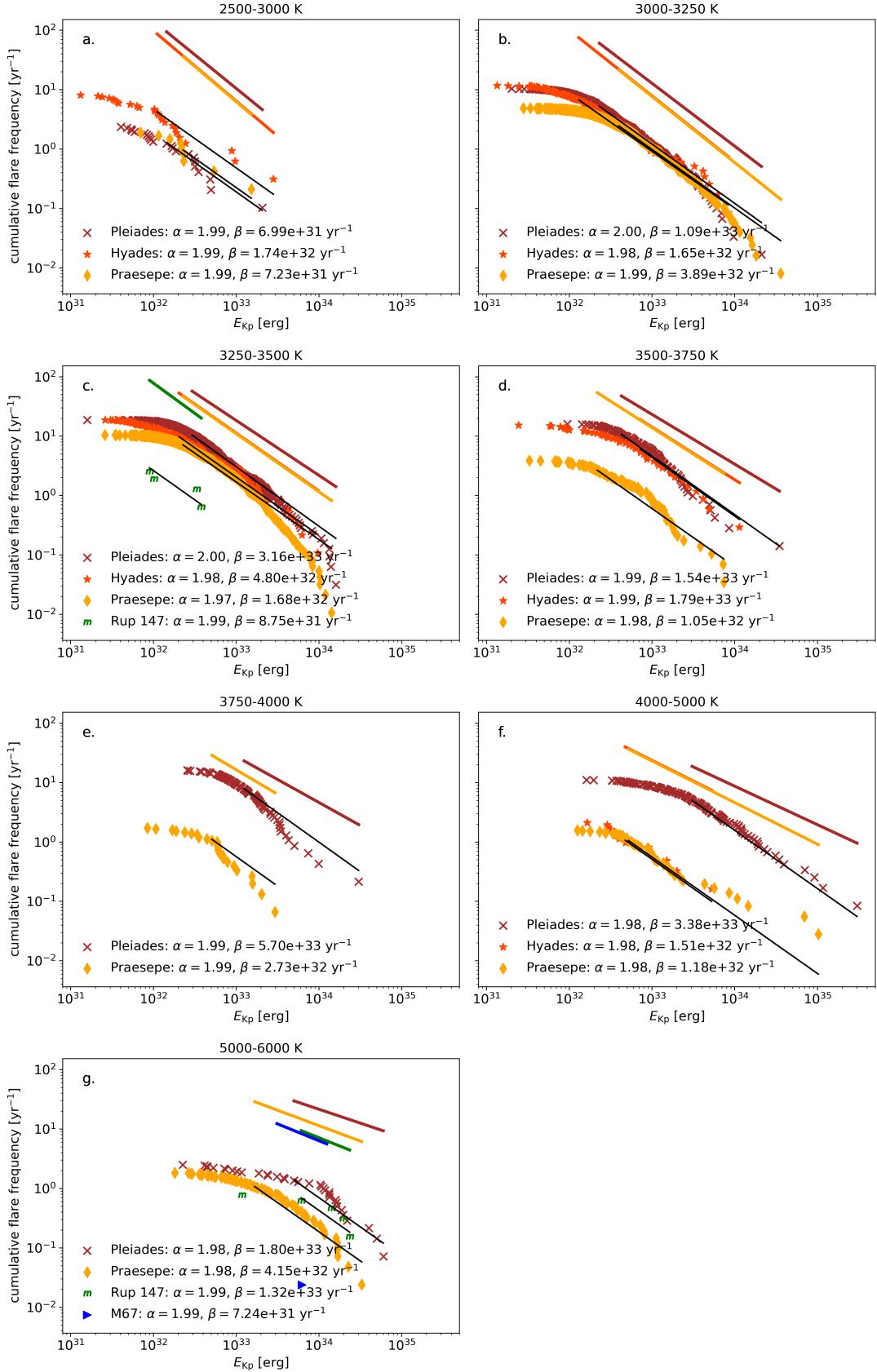
We determined  $L_{\text{bol},*}$  from  $R_*$  and  $T_{\text{eff}}$ , as described in Sec. 2.3. The energy released in flares was inferred using our derived stellar luminosities. It declines with age for every  $T_{\text{eff}}$  bin considered for both the total luminosity and relative to the quiescent flux (Fig. 8). However, we found that  $FA$  remained high ( $> 10^{-5}$ ) in early to mid M dwarfs for several hundred Myr

after Pleiades age.

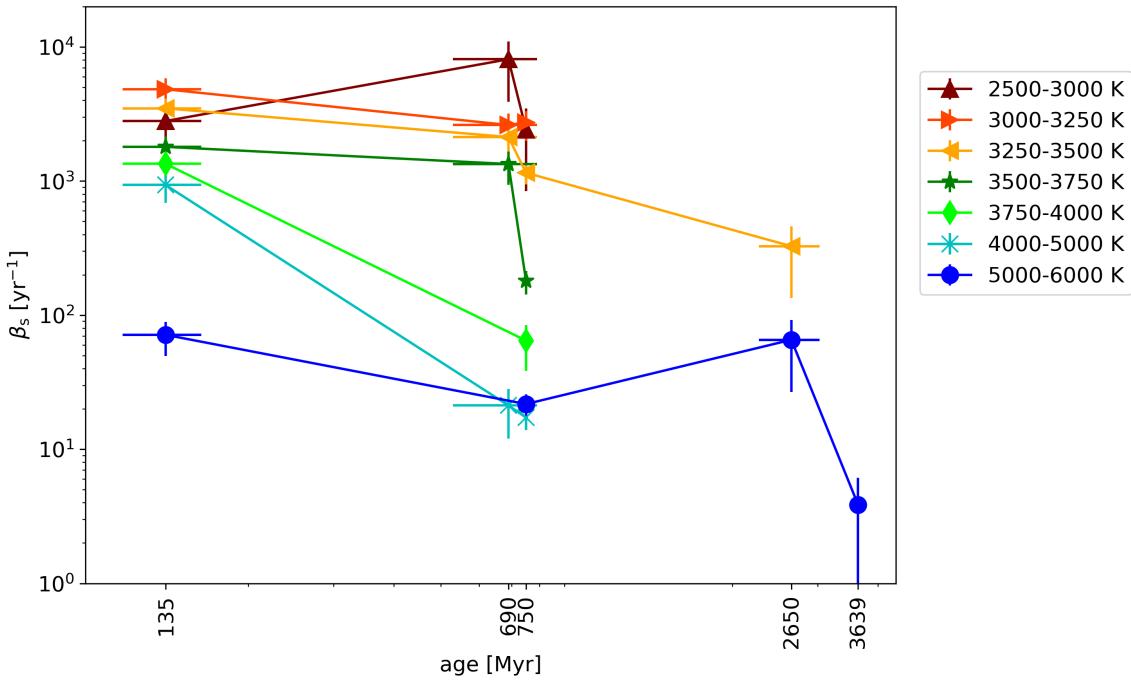
Lurie et al. (2015) analysed the flares on GJ 1245 A and B, two M5 dwarfs in a triple system with another M8 dwarf, that were observed in 1 minute cadence during the original Kepler mission. GJ 1245 A and B both fall roughly into the 3000 – 3250 K bin, and rotate at periode  $< 1$  d. The authors calculated  $FA \approx 10^{-4}$



**Fig. 5.** Cumulative flare frequency distributions (FFD, scatter) in  $ED$ , and respective power law fits (black lines), in log-log representation. In each of the panels a.-g. the distributions of flares are shown for different open clusters for samples of stars in the  $T_{\text{eff}}$  bin that is given in the panel title. The different clusters are the Pleiades (maroon crosses), the Hyades (orange stars), Praesepe (yellow diamonds), Ruprecht 147 (green "m"s), and M67 (blue triangles). The legends denote the best fit power law exponent  $\alpha$ , and intercept  $\beta$  (Eq. 3) for each distribution. We used the fit to the full sample as a prior for  $\alpha$ . Higher  $\beta$  indicate higher flare frequencies. The extension of the black lines in  $ED$  indicates the low-energy threshold adopted for every fit. For the single flare detection in M67 we give the power law parameters obtained directly from Eq. 3 using the prior value for  $\alpha$ . When no distribution is shown, either no flares were found in the sub-sample, or no stars in that  $T_{\text{eff}}$  bin were observed by K2. The overview of all power law fit parameters including uncertainties is given in Table 4.



**Fig. 6.** Same as in Fig. 5 but for  $E_{\text{Kp}}$  instead of  $ED$ . Additionally, predictions from the Davenport et al. (2019) gyrochronology-calibrated FFD parametrization with stellar age and mass were added in the colors that denote the different open clusters (see discussion in Sec. 5.2). To calculate the predictions, we used the mean of each  $T_{\text{eff}}$  bin to determine the mass (Pecaut & Mamajek 2013), and open cluster ages from the literature.



**Fig. 7.**  $\beta_s$  as a function of age grouped by  $T_{\text{eff}}$  (different colors and symbols in the legend) in log-log representation.  $\beta_s$  indicates the occurrence rate of flares with  $ED = 1$  s (see Eq. 3). It is controlled for stellar luminosity, so that we can compare flaring activity across both age, and  $T_{\text{eff}}$ , at the same time. The age labels correspond to the open cluster ages in Table 1.

**Table 3.** FFD and power law fit parameters to the full flare sample in  $E_{\text{Kp}}$  and  $ED$  (Fig. 4).

	$ED$	$E_{\text{Kp}}$
$\alpha$	$2.10(0.04)$	$1.98(0.03)$
$\beta$ [yr\$^{-1}\$]	$2023(600)_{460}$	$4.01(32.15)_{3.55} \cdot 10^{32}$
$n_{\text{tot}}$	3844	3844
$n_{\text{fit}}$	961	1213
min.	629 s	$6.63 \cdot 10^{32}$ erg
max.	82185 s	$3.01 \cdot 10^{35}$ erg

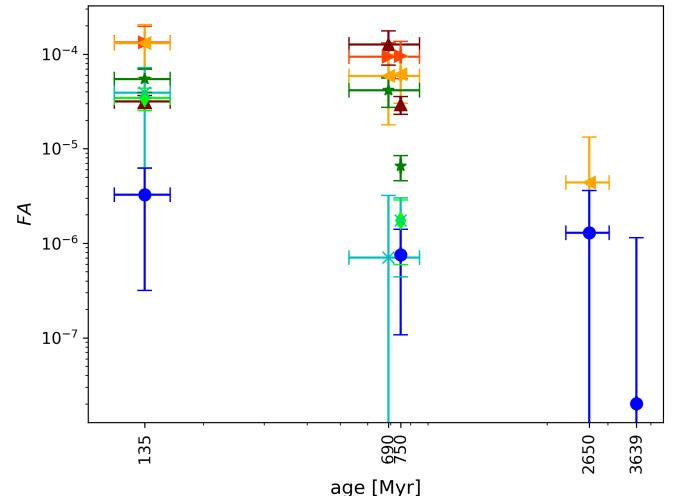
**Notes.** Uncertainties are given in parentheses. min. and max. are minimum and maximum detected flare  $ED$ /energy,  $n_{\text{tot}}$  is the total number of flares in the sample.  $n_{\text{fit}}$  is the number of flares used for the power law fit.

for both components, consistent with the  $FA$  value in the 135 Myr old Pleiades in this  $T_{\text{eff}}$ . But as the activity remains high at least up to the ages of Hyades and Praesepe, GJ 1245 could also be older.

#### 4.4. Hyades and Praesepe: Rotating differently

We expected the flaring activity in the Hyades and Praesepe to be very similar because they are nearly coeval clusters with ages around 700 Myr, and comparable metallicities ([Fe/H](Praesepe) = 0.16, [Fe/H](Hyades) = 0.13, Netopil et al. 2016). We found this reflected in all our FFDs, except for the 3500 – 3750 K temperature bin, where the Hyades appeared almost as active as the Pleiades (135 Myr), while Praesepe showed flare frequencies about ten times lower than the Pleiades. This discrepancy could be explained by the rotation period distribution in that bin.

More energetic flares can be expected from faster rotating



**Fig. 8.** Median  $FA$  and standard deviation plotted against age grouped by  $T_{\text{eff}}$  in log-log representation. The colors and shapes are the same as in Fig. 7. We used the shared energy threshold ( $3.31 \cdot 10^{32}$  erg) that we determined for the full FFD in  $E_{\text{Kp}}$  in Fig. 4 to calculate  $FA$  using Eq. 8.  $FA$  is a meaningful measure of relative stellar activity as long as the flux portion of the quiescent star in the Kepler band is roughly constant. It is therefore more meaningful to compare  $FA$  across age than across  $T_{\text{eff}}$ .

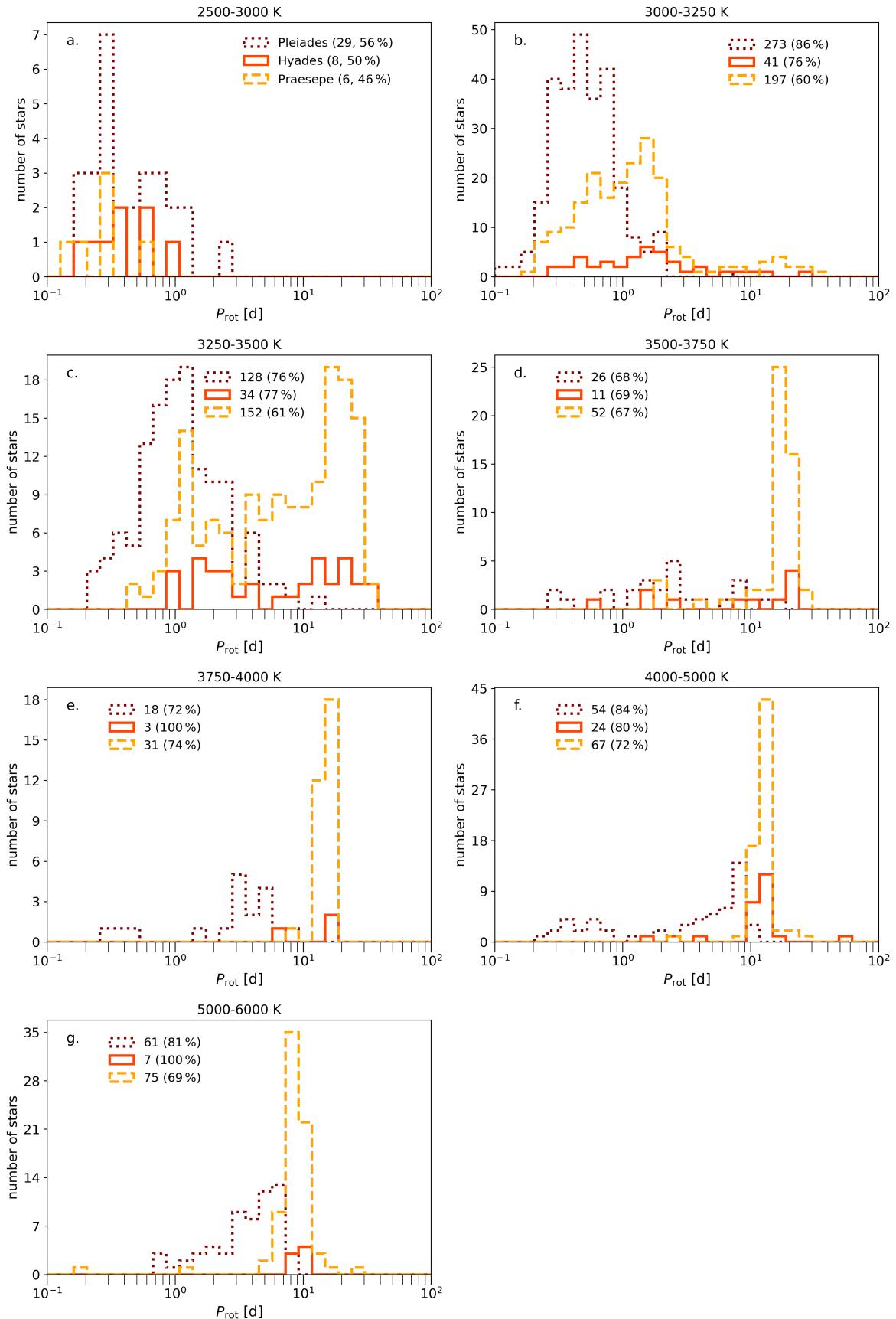
stars (Candelaresi et al. 2014; Doorsselaere et al. 2017; Yang et al. 2017). However, the famous age\$^{-1/2}\$ Skumanich spin-down law for rotation braking and Ca II depletion as a proxy to magnetic activity (Skumanich 1972) was only ever true to a first approximation.

One example of non-linear spin-down was the apparent temporary stalling of spin-down seen in K dwarfs in the 1 Gyr

**Table 4.** Power law fit parameters  $\alpha$  and  $\beta$  for all clusters and  $T_{\text{eff}}$  bins in  $E_{\text{Kp}}$  and  $ED$  FFDs.

$T_{\text{eff}}$ [K]	cluster	$\alpha_s$	$\beta_s$ [yr $^{-1}$ ]	$ED_{\text{min}}$ [s]	$ED_{\text{max}}$ [s]	$n_{\text{tot},s}$	$n_{\text{fit},s}$	$\alpha_{\text{erg}}$	$\beta_{\text{erg}}$ [yr $^{-1}$ ]	$E_{\text{Kp,min}}$ [erg]	$E_{\text{Kp,max}}$ [erg]	$n_{\text{tot,erg}}$	$n_{\text{fit,erg}}$
2500-3000	Hyades	2.10 (0.03)	8102 (4201) (2853)	1.24 · 10 <sup>3</sup>	1.66 · 10 <sup>4</sup>	26	9	1.99 (0.03)	1.7 (1.5) · 10 <sup>32</sup>	1.11 · 10 <sup>32</sup>	2.82 · 10 <sup>33</sup>	26	13
	Pleiades	2.10 (0.03)	2803 (1290) (916)	0.99 · 10 <sup>3</sup>	1.52 · 10 <sup>4</sup>	23	12	1.99 (0.03)	0.7 (0.6) · 10 <sup>32</sup>	1.45 · 10 <sup>32</sup>	2.08 · 10 <sup>33</sup>	23	12
3000-3250	Præsepe	2.10 (0.03)	2427 (1588) (1043)	1.13 · 10 <sup>3</sup>	0.86 · 10 <sup>4</sup>	9	4	1.99 (0.03)	0.7 (5.7) · 10 <sup>32</sup>	2.3 · 10 <sup>32</sup>	1.53 · 10 <sup>33</sup>	9	4
	Hyades	2.08 (0.03)	2618 (709) (570)	2.37 · 10 <sup>2</sup>	1.29 · 10 <sup>4</sup>	137	76	1.98 (0.03)	1.6 (11.4) · 10 <sup>32</sup>	1.31 · 10 <sup>32</sup>	0.84 · 10 <sup>34</sup>	137	78
3250-3500	Pleiades	2.09 (0.03)	4842 (1262) (1011)	0.58 · 10 <sup>3</sup>	0.63 · 10 <sup>5</sup>	616	258	2.00 (0.03)	1.1 (6.3) · 10 <sup>33</sup>	2.32 · 10 <sup>32</sup>	2.11 · 10 <sup>34</sup>	616	311
	Præsepe	2.08 (0.03)	2739 (743) (578)	0.7 · 10 <sup>3</sup>	0.82 · 10 <sup>5</sup>	601	275	1.99 (0.03)	0.4 (2.3) · 10 <sup>33</sup>	0.4 · 10 <sup>33</sup>	0.36 · 10 <sup>35</sup>	601	302
3500-3750	Hyades	2.09 (0.03)	2130 (542) (426)	1.37 · 10 <sup>2</sup>	0.36 · 10 <sup>4</sup>	174	83	1.98 (0.03)	0.5 (3.2) · 10 <sup>33</sup>	2.04 · 10 <sup>32</sup>	0.97 · 10 <sup>34</sup>	174	90
	Pleiades	2.08 (0.03)	3487 (77) (63)	2.19 · 10 <sup>2</sup>	1.62 · 10 <sup>4</sup>	590	310	2.00 (0.03)	3.2 (18.0) · 10 <sup>33</sup>	2.95 · 10 <sup>32</sup>	1.61 · 10 <sup>34</sup>	590	327
3750-4000	Præsepe	2.02 (0.03)	1153 (211) (173)	1.59 · 10 <sup>2</sup>	1.03 · 10 <sup>4</sup>	970	612	1.97 (0.02)	1.7 (7.5) · 10 <sup>32</sup>	2.28 · 10 <sup>32</sup>	1.42 · 10 <sup>34</sup>	970	663
	Ruprecht 147	2.10 (0.03)	326 (193) (132)	0.67 · 10 <sup>2</sup>	2.81 · 10 <sup>2</sup>	4	4	1.99 (0.03)	0.9 (6.9) · 10 <sup>32</sup>	0.9 · 10 <sup>32</sup>	0.38 · 10 <sup>33</sup>	4	4
4000-5000	Hyades	2.10 (0.03)	1340 (400) (315)	0.93 · 10 <sup>2</sup>	2.88 · 10 <sup>3</sup>	52	28	1.99 (0.03)	1.8 (13.5) · 10 <sup>33</sup>	0.72 · 10 <sup>33</sup>	1.14 · 10 <sup>34</sup>	52	20
	Pleiades	2.10 (0.03)	1800 (440) (357)	1.06 · 10 <sup>2</sup>	0.57 · 10 <sup>4</sup>	113	69	1.99 (0.03)	1.5 (0.8) · 10 <sup>33</sup>	0.43 · 10 <sup>33</sup>	0.35 · 10 <sup>35</sup>	113	75
5000-6000	Præsepe	2.08 (0.03)	180 (38) (32)	0.46 · 10 <sup>2</sup>	1.63 · 10 <sup>3</sup>	111	74	1.98 (0.03)	1.1 (7.5) · 10 <sup>32</sup>	2.19 · 10 <sup>32</sup>	0.74 · 10 <sup>34</sup>	111	76
	Pleiades	2.11 (0.03)	1346 (380) (300)	0.98 · 10 <sup>2</sup>	2.05 · 10 <sup>3</sup>	75	35	1.99 (0.03)	0.6 (4.5) · 10 <sup>34</sup>	1.25 · 10 <sup>33</sup>	3.03 · 10 <sup>34</sup>	75	36
6000-7000	Præsepe	2.10 (0.03)	64 (26) (19)	0.56 · 10 <sup>2</sup>	0.37 · 10 <sup>3</sup>	26	10	1.99 (0.03)	2.7 (2.4) · 10 <sup>32</sup>	0.51 · 10 <sup>33</sup>	2.95 · 10 <sup>33</sup>	26	16
	Hyades	2.10 (0.03)	21 (6) (6)	1.22 · 10 <sup>1</sup>	1.1 · 10 <sup>2</sup>	13	7	1.98 (0.03)	1.5 (1.4) · 10 <sup>32</sup>	0.48 · 10 <sup>33</sup>	0.53 · 10 <sup>34</sup>	13	6
7000-8000	Pleiades	2.09 (0.03)	935 (246) (97)	1.17 · 10 <sup>2</sup>	0.46 · 10 <sup>4</sup>	131	55	1.98 (0.03)	0.3 (2.6) · 10 <sup>34</sup>	3.06 · 10 <sup>33</sup>	3.01 · 10 <sup>35</sup>	131	59
	Præsepe	2.08 (0.03)	17 (2) (2)	0.99 · 10 <sup>1</sup>	1.52 · 10 <sup>3</sup>	56	47	1.98 (0.03)	1.2 (8.6) · 10 <sup>32</sup>	0.52 · 10 <sup>33</sup>	1.03 · 10 <sup>35</sup>	56	37
8000-9000	M67	2.10 (0.03)	3 (3) (2)	0.6 · 10 <sup>2</sup>	0.6 · 10 <sup>2</sup>	1	1	1.99 (0.03)	0.7 (0.7) · 10 <sup>32</sup>	0.62 · 10 <sup>34</sup>	0.62 · 10 <sup>34</sup>	1	1
	Pleiades	2.09 (0.03)	71 (21) (17)	0.34 · 10 <sup>2</sup>	0.6 · 10 <sup>3</sup>	35	19	1.98 (0.03)	1.8 (4.8) · 10 <sup>33</sup>	0.5 · 10 <sup>34</sup>	0.6 · 10 <sup>35</sup>	35	19
9000-10000	Præsepe	2.10 (0.03)	21 (4) (3)	1.49 · 10 <sup>1</sup>	1.85 · 10 <sup>2</sup>	76	42	1.98 (0.03)	0.4 (3.1) · 10 <sup>33</sup>	1.68 · 10 <sup>33</sup>	0.33 · 10 <sup>35</sup>	76	44
	Ruprecht 147	2.10 (0.03)	65 (38) (26)	0.56 · 10 <sup>2</sup>	2.16 · 10 <sup>2</sup>	5	4	1.99 (0.03)	1.3 (11.1) · 10 <sup>33</sup>	0.61 · 10 <sup>34</sup>	2.36 · 10 <sup>34</sup>	5	4

**Notes.** Uncertainties are given in parentheses. Indices <sub>s</sub> and <sub>erg</sub> refer to the units of  $ED$  and  $E_{\text{Kp}}$  FFDs.  $E_{\text{Kp,min}}/ED_{\text{min}}$  and  $E_{\text{Kp,max}}/ED_{\text{max}}$  are minimum and maximum detected flare  $ED$ /energy,  $n_{\text{tot}}$  is the total number of flares in the sample.  $n_{\text{fit}}$  is the number of flares used for the power law fit.



**Fig. 9.** Rotation histograms for the cluster members that were searched for flares and had measured rotation periods  $P_{\text{rot}}$ . Rotation periods were taken from Douglas et al. (2019) for the Hyades (orange solid) and Praesepe (yellow dashed), and from Rebull et al. (2016) for the Pleiades (maroon dotted). The numbers and percentages in the legends in each panel denote the number of stars for which periods were given, and what fraction of the full sample that represented, respectively. The period distributions of Praesepe and the Hyades are consistent in all  $T_{\text{eff}}$  bins but the 3500 – 3750 K bin, where the distribution of the Hyades appears more similar to the younger Pleiades than the same-age Praesepe.

**Table 5.** Possible binary configurations for flaring old stars.

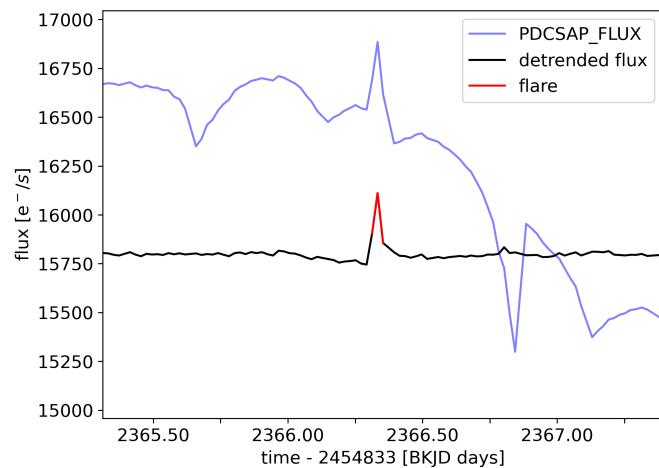
cluster	EPIC	median SpT	binary
M67	211434440	K1	K2 + M5.5
Ruprecht 147	219601739	G8	K1 + M6
Ruprecht 147	219610232	K0.5	K2 + M5.5
Ruprecht 147	219591752	M3	M3.5 + M3.5

old open cluster NGC 6811 (Curtis et al. 2019). As our results indicate, the decline of flaring activity with age did not follow the Skumanich law across the stellar main sequence life time. Moreover, it depended crucially on  $T_{\text{eff}}$  (see Fig. 7).

In Fig. 9 we show rotation periods derived from K2 light curves for the Pleiades (Rebull et al. 2016), the Hyades and Praesepe (Douglas et al. 2019), to illustrate the  $P_{\text{rot}}$  distributions that correspond to our  $T_{\text{eff}}$  bins.

In all but the 3500 – 3750 K temperature bin the rotation period distributions of the Hyades and Praesepe were similar, while the Pleiades showed shorter periods on average. Only in this bin the rotation periods found in the Hyades and the Pleiades were more alike than in the Hyades and Praesepe. In this temperature regime, the majority of Praesepe stars had rotation periods  $> 10$  d, while in the Hyades and the Pleiades rotation periods were more evenly distributed in the 0.3 – 30 d range. So the rotation period distribution made a jump between the ages of Hyades and Praesepe, and the flaring rates echoed this decrease. This may be the time-resolved effect of rapid transition in stellar activity previously detected in various observables, including flaring luminosities and amplitudes, at the  $P_{\text{rot}} = 10$  d boundary (Stelzer et al. 2016; Lu et al. 2019).

#### 4.5. M67 and Ruprecht 147: Possibly binaries



**Fig. 10.** Single confirmed flare candidate in M67 on the K1 dwarf EPIC 211434440 (red line). Blue: Pre-search Data Conditioning (PDCSAP) flux, that is, the light curve before K2SC de-trending. Black: De-trended light curve that was searched for flares with AltaiPony. The amplitude of the flare is 2 %, and the energy emitted in the Kepler band is  $E_{\text{Kp}} = 6.25 \times 10^{33}$  erg .

We found 4 and 8 flaring members in M67 and Ruprecht 147, respectively. Upon close inspection the majority were false positives, or occurred on multiple systems or evolved stars that were not properly filtered out. Most flare candidates in these old clusters were detected on RS CVn binaries, cataclysmic binaries,

spectroscopic binaries, and red giant stars. Excluding all these, we were left with one flare in M67 on an early K dwarf (Fig. 10). In Ruprecht 147, we excluded EPIC 219426848, a double-lined spectroscopic binary (Curtis et al. 2013), and so narrowed down the list to a flare on a G8 star in Ruprecht 147, and four flares each on a K0.5 and an M3 star. For these stars, the multiplicity status was unknown. The uncertainty about the stellar properties suggests a range of possible masses of the system. The mass budget could be calculated from the uncertainties on their radii  $R_*$  using the relations from Eker et al. (2018). In all cases, it was large enough that the stars in question could in principle be binary stars with mid-M dwarf companions that were too faint to be detected. We calculated hypothetic binary pairs for the cases where the primary mass was the smallest possible within  $1\sigma$  on  $R_*$ . We give the median spectral type of the target if it was a single star, and the spectral types of the possible binaries in Table 5.

After the manual inspection of the younger clusters we concluded that, while binaries and false positives were present in these clusters also, they were not the dominating source of flares. As multiplicity rates for low mass stars decrease from about 50% for F type stars (Raghavan et al. 2010) to 22% for L and T dwarfs (Duchêne & Kraus 2013), with M dwarfs showing somewhat higher multiplicity rates at about 27% (Winters et al. 2019), we expect that the error introduced by unresolved binaries on our results was smaller for the cooler  $T_{\text{eff}}$  bins.

## 5. Discussion

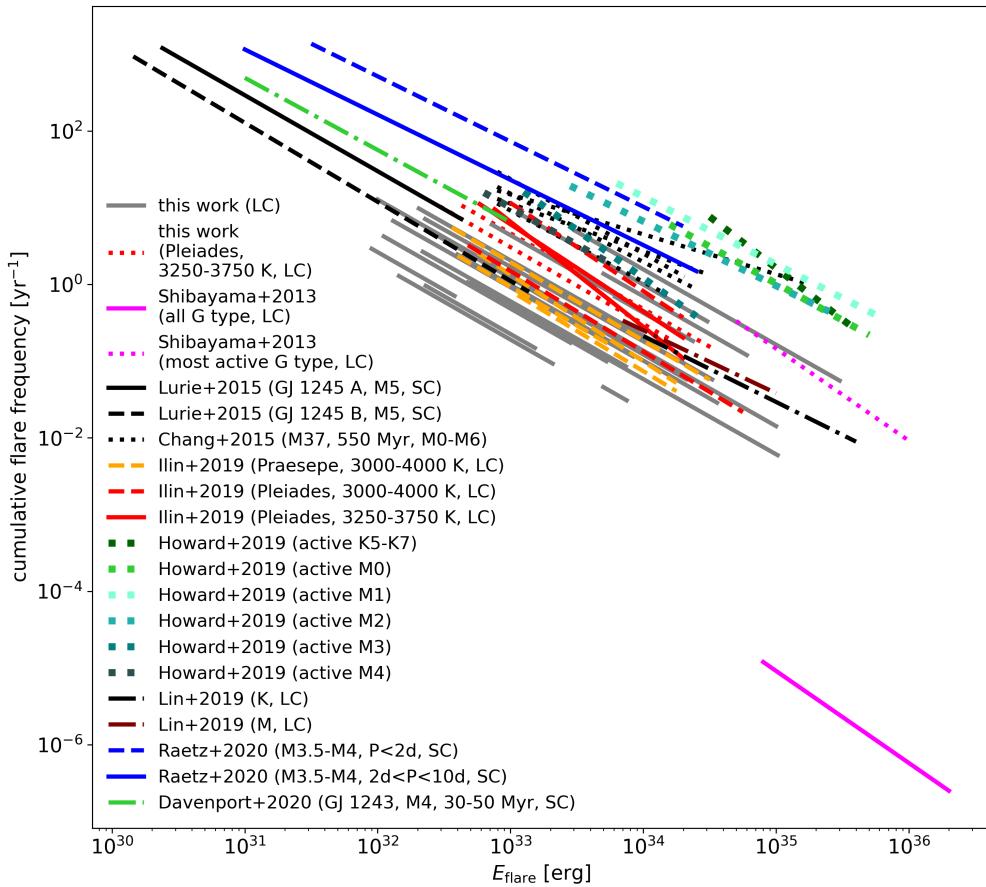
The launch of the Kepler satellite in 2009 caused a surge in statistical flare studies, a trend that the Transiting Exoplanet Survey Satellite will continue and expand. Comparing our results with these works (Shibayama et al. 2013; Lurie et al. 2015; Lin et al. 2019; Raetz et al. 2020) and recent ground based surveys (Chang et al. 2015; Howard et al. 2019) we found encouraging consistency but also some noteworthy discrepancies that we could not always trace back to its cause (Section 5.1). We discuss the disagreement between our FFDs and a gyrochronological model recently proposed for its parametrization as a function of mass and age (Section 5.2, Davenport et al. 2019). We conclude the discussion with a reflection on the observational limitations of the representation of FFDs by power law distributions (Section 5.3), and the possibility of observing the maximum possible flare energy in our FFDs (Section 5.4).

### 5.1. Consistency with statistical flares studies

We found our results to be broadly consistent with previous work both on Kepler/K2 data, and ground based surveys in the optical regime, with some notable exceptions. Fig. 11 presents the FFDs of the studies that we discuss in this context. Overall we believe that the differences between individual results were rooted in a combination of astrophysical properties like age, mass, rotation speed, and the influence of selection criteria on the respective samples.

The flare frequency distributions of two fast rotating (0.26 and 0.71 days) M5 flare stars in a M5-M5-M8 triple system (Lurie et al. 2015) were consistent with the FFDs of flares in the Pleiades, Praesepe, and Hyades in the 3000 – 3250 K bin.

We also found that the FFDs of nearby K and M dwarfs, observed by Kepler/K2 within 200pc and 100pc, respectively, fell in the range between the Pleiades and Hyades/Praesepe FFDs in our sample (Lin et al. 2019).



**Fig. 11.** Comparison of FFDs found in different flare studies. Grey lines: Flare frequency distributions from this work. The one-flare sample for 5000–6000 K in M67 is indicated as the shortest visible grey line. Red solid lines, and red and orange dashes: K2 long cadence light curves study of the Pleiades and Praesepe, respectively, in four  $T_{\text{eff}}$  bins from 3000 to 4000 K each (Paper I). Magenta line and dots: Superflares on all G dwarfs in Kepler, and its most active sub-sample (Shibayama et al. 2013). Black line and dashes: Two M5 stars in an M5-M5-M8 triple system observed by Kepler (Lurie et al. 2015). Black dots: Flare study based on MMT survey of the  $\sim 550$  Myr old open cluster M37 (Chang et al. 2015). Dots in shades of green to blue: Evryscope all-sky flare search on late K to mid M dwarfs (Howard et al. 2019). Black and brown dash-dotted lines: K2 long-cadence study of K and M dwarfs within 200pc and 100pc (Lin et al. 2019). Dark blue line and dashes: K2 short cadence light curves of M3.5–M4 dwarfs with known rotation periods below 2 d and between 2 and 10 d respectively (Raetz et al. 2020). Green dash-dotted line: The 30 – 50 Myr old M4 flare star GJ 1243 exhibited a constant flaring activity for over a decade of intermittent space based observations with Kepler and TESS (Davenport et al. 2020). In the studies that are based on Kepler or K2 observations, "SC" and "LC" denote the use of 1 minute and 30 minute cadence time series, respectively.

The frequency of superflares at  $10^{34}$  erg in the most active G-dwarfs and Sun-like stars in Kepler was once in 10–100 days, and every 440 years on average solar-type stars (Shibayama et al. 2013). Benchmarking on the flare detected on a K1 dwarf in M67 in the G to early K bin (5000–6000 K), and accounting for the fraction of the total flare energy that the Kepler passband covers at a flare temperature of 10 000 K, we found the flaring rate at this energy to amount to about once every 100 years. Doing the same for the Pleiades stars yielded that ZAMS G to K dwarfs flare once per year at  $10^{34}$  erg. Our young G-K stars therefore were less active than the active G dwarfs in Shibayama et al. (2013), while our solar-age G-K dwarfs appeared more active than in their results.

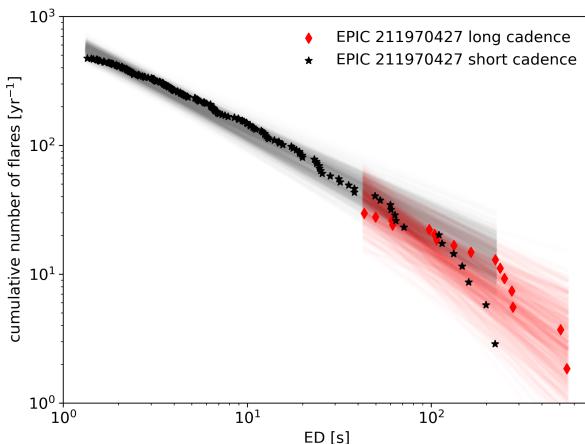
Howard et al. (2019) monitored superflares on cool stars with bolometric energies above  $10^{33}$  erg and up to  $10^{36}$  erg. They found power law exponents  $\sim 2$  in FFDs resolved by spectral type. The activity levels for spectral types K5 to M4 were on average higher than in our sample. However, Howard et al. (2019) only included stars that exhibited flares (active stars) into their FFDs, whereas we also accounted for the observing time of

stars where no flares had been detected.

Photometric flares observed by the MMT 6.5 telescope (Hartman et al. 2008) in the  $\sim 550$  Myr old open cluster M37 (Chang et al. 2015) appeared on average more active than the Pleiades in our study for spectral types M0–M6. This could imply that flaring activity in these low mass stars peaked not at Pleiades age but at a later evolutionary stage.

Similarly, Raetz et al. (2020) pointed out that their fast rotating M1.5–M4.5 dwarfs appeared more active than the results for the Pleiades in Paper I in the 3250 – 3750 K range. Our revised results remained consistent with Paper I (Fig. 11). Both samples were comprised of flares with periods below 10 d. Rotation periods  $P_{\text{rot}}$  were given for 100% of the sample in Raetz et al. (2020), and for > 68% in our Pleiades sub-sample (Fig. 9).

We consider three plausible reasons for the discrepancy here. First, the M1.5–M4.5 dwarfs in Raetz et al. (2020) could be at an age at which stars show higher flaring activity than stars with the same rotation periods in the Pleiades. Second, there may be selection effects in the short cadence sample in Raetz et al. (2020) as compared to our sample. Targets for short



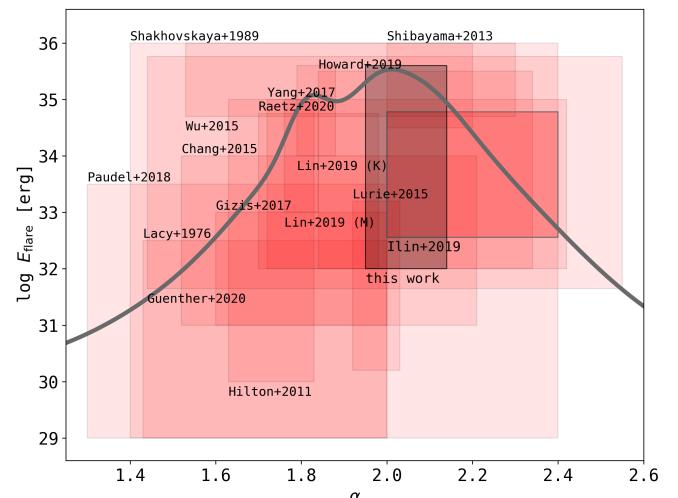
**Fig. 12.** FFDs in  $ED$  for EPIC 211970427, an M3 dwarf in Praesepe that was observed in both short 1 minute cadence (campaigns 5, 18) and long 30 minute cadence (campaigns 5, 16, and 18). We detrended, and searched both light curve types for flares using the method described in Section 3.1. Red diamonds indicate flares found in the long cadence light curves, black stars are flares from the short cadence light curves. The red and black lines were randomly sampled from the posterior distributions of the power law fits to the long and short cadence FFDs, respectively (see methods in Section 3.4).

cadence observations in K2 were not selected randomly from the underlying stellar population, but were filtered for a variety of properties that biased the final selection towards more active stars. Third, our  $3250 - 3750$  K bin corresponds rather to spectral types M1-M3.5 than to M1.5-M4.5 (Pecaut & Mamajek 2013). So the  $T_{\text{eff}}$  bins chosen in Raetz et al. (2020) encompassed later, and therefore more active, spectral types. Eventually, a resolution of this discrepancy must also take into account that the results for our Praesepe sample in the  $3250 - 3750$  K range were consistent with the slow rotators ( $P_{\text{rot}} \sim 32$  d) in Raetz et al. (2020).

Finally, we note that the difference in time sampling between our study (30 minute cadence), and 1 minute cadence in Raetz et al. (2020) can most likely be ruled out as an explanation for the discrepancy. In Fig. 12, we show an example that illustrates that short cadence light curves used in Raetz et al. (2020), and long cadence light curves studied here yielded FFDs that were consistent withing uncertainties. There were no exactly matching events in the FFDs although a subset of the short cadence flares was detected in long cadence also for two main reasons. First, in long cadence, flare energy was typically underestimated (Yang et al. 2018). Second, smaller flares were merged together if they occurred within a sufficiently short time period. What looked like one flare in long cadence was sometimes resolved into two or more flares in short cadence.

## 5.2. Gyrochronological FFD parametrization

Davenport et al. (2019) derived a parametrization of FFD parameters  $\alpha$  and  $\beta$  as a function of mass and age using the Kepler flare catalog (Davenport 2016), rotation periods from McQuillan et al. (2014), and the Mamajek & Hillenbrand (2008) gyrochrone model. We show the model power laws in Fig. 6 alongside our results. While the power law exponents appeared consistent with our estimates, the activity level in our sample was much lower than predicted by the model. A deviation



**Fig. 13.** Literature overview over power law fits to FFDs. Red rectangles indicate the energy range in which flares were detected (height), and the range of values determined for the power law exponent  $\alpha$  (width) including quoted uncertainties and superimposing the ranges for  $\alpha$  from multiple FFDs if any. The references are anchored to the upper or lower left corner of the respective rectangle. The grey rectangle refers to this study. The grey line depicts the superposition of these results, representing each by a Gaussian fit with the range used as standard deviation, and weighted by the square root of the number of flares found in each study. Paper I is shown with a dark outline. See Section 5.3 for details on the quoted works.

was expected as the model was not designed to predict absolute ages. Davenport et al. (2019) noted a sample bias in their study towards more active stars. As discussed by the authors, their model overpredicted the superflaring rate of the average Sun-like sample from Shibayama et al. (2013) and more resembled the rate for their most active sub-sample. Moreover, the model was not designed to be applicable to flaring activity in the saturated regime of fast young rotators (Jackman et al. 2020). With our data we quantify the overestimation of flaring rates to be one to two orders of magnitude. Our results also suggest that the model overpredicts the activity level for old stars more than for younger stars (Fig. 6).

## 5.3. Universality of the power law exponent $\alpha$

Taking into account uncertainties and systematic errors resulting from the use of different power law fitting methods (Maschberger & Kroupa 2009), the power law exponent  $\alpha \sim 2$  appears to be similar for the majority of studies on flare statistics in the literature so far, irrespective of spectral type or age. To illustrate this, we searched the literature for studies where power laws were fitted to FFDs of flares in the optical regime using different methods (Fig. 13).

Lacy et al. (1976) pioneered in statistical flare studies using observations of eight UV Ceti stars. Shakhevskaya (1989) analyzed monitoring data from several dozen active early K to late M dwarfs in the solar vicinity, the Orion, the Pleiades, the Hyades, and the Praesepe cluster. A number of studies focused on solar-type and G dwarfs stars in Kepler (Shibayama et al. 2013; Wu et al. 2015). Hilton (2011) derived  $\alpha$  from ground-based observations of M dwarfs. Lurie et al. (2015) determined  $\alpha$  from Kepler short cadence light curves of two active M5 dwarfs in Kepler. Yang et al. (2017) followed up on Davenport (2016)

in searching the entire Kepler catalog for flares. Lin et al. (2019) (M) and (K) were based on nearby M and K dwarf flares in K2 long cadence light curves. (Raetz et al. 2020) studied short cadence M dwarfs with rotation periods in K2. Gizis (2017) constructed an FFD for 22 flares on an M7 dwarf. Paudel et al. (2018) studied M6-L0 dwarfs in K2 short cadence data. Howard et al. (2019) found over 500 superflares on 284 stars in the framework of the Evryscope all-sky survey. Photometric flares were observed by the MMT 6.5 telescope (Hartman et al. 2008) in the  $\sim 550$  Myr old M37 open cluster (Chang et al. 2015). Finally, (Günther et al. 2020) took a first look to  $\alpha$  as derived from stars observed during TESS Cycle 1. From their work we only use  $\alpha$  derived from stars with 20 or more flare detections in Fig. 13.

While most studies consistently found values somewhat above or below  $\alpha \approx 2$ , the comparison would reveal unresolved systematic errors in all these studies, including our own, if  $\alpha$  was truly universal for all FFDs. The values derived for  $\alpha$  in our sample were close to 2 in  $ED$  space, and  $\sim 2.1$  in  $E_{kp}$ , consistent with previous work. The determination of  $\alpha$  using different methods without reliable uncertainty estimates across different studies makes it difficult to assess whether the spread in  $\alpha$  from about 1.4 to 2.5 is physically motivated or not. Additionally, for both low and high flare energies, the detection of flares is usually incomplete or contaminated. This is caused by finite time sampling, cosmic ray impact, and scatter in the light curve at the low end, and finite observation time at the high end (see Section 3.5). As long as these methodological and observational caveats apply, the question about the universal nature of  $\alpha$  cannot be settled.

#### 5.4. Maximum flare energy

Spots can survive on the stellar surface from a few days to nearly a year (Namekata et al. 2017; Davenport 2015). Complex spot geometry is correlated with the strongest X-class flares on the Sun (Toriumi et al. 2017; Sammis et al. 2000). These observations support the idea that flares are associated with the presence of certain types of starspots, or more generally, certain types of active regions. Since we expect that there is a maximum flare energy a spot can produce, the underlying power law relation must break a some  $ED_{\max}$  in any FFD. We tested a possible truncation of our FFDs using the Kolmogorov-Smirnov test for consistency with a power law distribution, followed by the exceedance statistic to test for a truncation of this power law as suggested by Maschberger & Kroupa (2009). While most FFDs with  $\geq 50$  flares were consistent with a power law distribution, we found no evidence for truncation in any of the  $T_{\text{eff}}$ -age bins. This would imply that we did not sample the highest possible energies. Such a conclusion is plausible because much higher relative fluxes than our observed maximum flare amplitudes were detected on other stars (Paudel et al. 2018; Jackman et al. 2019; Schmidt et al. 2016).

## 6. Summary and Conclusions

In this study, we investigated how flaring activity unfolded for different age and  $T_{\text{eff}}$  ranges anchoring the age to the membership in five open clusters. We found flaring activity to decrease from mid M dwarfs to G stars, and from ZAMS to solar age, except for stars later than M5, where activity increased from ZAMS to 690 Myr.

Using multiple cluster membership studies we selected for high-probability members, and drew from a host of multiband

photometry catalogs and Gaia astrometry to determine  $T_{\text{eff}}$  and stellar luminosities. Keeping only targets with well-determined mass and age we proceeded to search 3435 K2 long cadence light curves obtained for these stars throughout the K2 mission. We applied the open source software K2SC to remove rotational variability and instrumental effects. We developed the open source software AltaiPony, and used it to automatically detect and determine the properties of flare candidates. We vetted all flare candidates manually and discussed various sources of false positives and incompleteness on both the high and the low energy end of the FFDs. All in all, we searched 2111 stars in five open clusters, and found 3844 flares. Most flares originated in the Pleiades and Praesepe, several hundred were found in the Hyades, a handful in Ruprecht 147, and only one flare candidate appeared on a K1 dwarf in M67.

We fitted power law functions to the flare frequency distributions of flares from stars in different age and  $T_{\text{eff}}$  bins. We found that the power law exponent  $\alpha$  was consistent with previous work, and caution against interpreting apparent trends, especially when comparing across different studies.

Our results indicated that flaring activity declined with age, and did so faster for higher mass stars. We provide the measured flaring rates in Table 4 as a contribution to the endeavor of mapping the flaring-age-mass relation from pre-main sequence to main sequence turn-off, and from the onset of an outer convection zone in F-type stars to the coolest brown dwarfs.

Our findings paint a picture of clear trends in the flaring evolution of GKM dwarfs on the main sequence. Regardless of varying methods employed for flare search and FFD analysis in previous studies we found encouraging consistency in flaring rates with both Kepler and K2 based work, and ground-based surveys. We noticed some differences, which we could not always resolve as being either systematic or astrophysical. We suggest that discrepancies between our results and flare studies that used rotation periods for their age estimates (Davenport et al. 2019; Raetz et al. 2020) could be explained by sample selection bias but may also point to limitations of rotation periods as an age indicator. Although some discrepancies in flaring activity between Praesepe and the Hyades could be traced back to their rotation period distributions in this study, the relation between flaring, rotation, and age awaits further investigation.

Although Kepler and K2 data were already used in more than 2400 publications in 2018, the public archive was still considered understudied (Barentsen et al. 2018). It will remain a benchmark mission not only for exoplanet search, but also for various subfields in stellar astronomy, and flaring activity research, in particular. Results from the first year of operations of TESS indicate that the light curve quality delivered by the mission so far was well suited for statistical flare studies, too (Doyle et al. 2020; Günther et al. 2020). The expected lifetime of TESS could be up to 20 years, in which case the number of high quality flare samples will soon outclass Kepler and K2 as the primary source of flares. The mission will vastly expand the treasury of light curves, leaving us optimistic that this work was only a first step towards a comprehensive flare rate evolution model.

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## Appendix A: Membership probabilities

To match catalogs on RA and declination we used the `astroML.crossmatch` tool for Python (Vanderplas et al. 2012). For the studies with classifiers we assigned membership probabilities as follows. In Gonzalez (2016):

$$\begin{aligned} p(M(\text{member})) &= 0.9, \\ p(BM(\text{binary member})) &= 0.9, \\ p(N(\text{non-member})) &= 0.1, \\ p(SN(\text{single non-member})) &= 0.1, \\ p(BN(\text{binary non-member})) &= 0.1, \\ p(U(\text{unknown member})) &= 0.5. \end{aligned}$$

In Curtis et al. (2013):

$$\begin{aligned} p(Y(\text{highest confidence member})) &= 0.9, \\ p(P(\text{possible/probable member})) &= 0.7, \\ p(N(\text{not likely/non-member})) &= 0.7, \\ p(B(\text{photometry consistent with blue stragglers})) &= 0.0. \end{aligned}$$

In Rebull et al. (2017):

$$\begin{aligned} p(\text{(best)}) &= 0.9, \\ p(\text{(ok)}) &= 0.6, \\ p(\text{(else)}) &= 0.1. \end{aligned}$$

Members from Rebull et al. (2016); Douglas et al. (2017); and Gaia Collaboration et al. (2018a) were assigned  $p = 0.9$  if they appeared in the final catalog.

Table A.1 gives an overview over different membership catalogs. Figure A.1 shows membership probability histograms of the final sample broken down by membership source.

## Appendix B: Cluster parameters

We compiled various age, distance, and metallicity values from the literature for the clusters in this study. Table B.1 shows that most studies independently arrived at similar results for all our clusters. The adopted parameters are emphasized, and were chosen based on consistency with prior work, uncertainties provided with the results and the data used for the study (newer catalogs were preferred).

## Appendix C: Solar System Object detection with SkyBot

K2 photometry was heavily affected by crossings of Solar System Objects (SSO), which consequently produced mimicking flare profiles in the light curves (Szabó et al. 2015). Independently, two of us visually verified the SSO-nature of 42 flaring light curve events by inspecting their corresponding TargetPixelFiles (TPF) using `lightkurve` (Lightkurve Collaboration et al. 2018). Except for one occasion, we found all events to be known SSOs, both in the visual survey and automatically using SkyBot (Berthier et al. 2016). Subsequently, we ran SkyBot on the remaining catalog of flares to remove other false positive events caused by known SSOs. Visual inspections of TPFs were carried out for any SSO detected within 30" of its target star at mid-event time. This yielded additional 103 contaminated signals distributed among three subgroups, (i) 77

unambiguous SSO-contaminants verified through `lightkurve` (ii) 18 faint SkyBot-detections ( $V \approx 20.4 - 23.8$ ) within  $\sim 3$  pixels from the respective target star but visually unresolvable in the TPF, (iii) 8 occurrences in which both flare and SSO-crossing occurred during the given flare duration stated in the catalog. The latter should be dealt with caution due to a high risk of inaccurate catalog information altered by the presence of SSOs. Despite of discarding a total of 144 flare events, we expect additional contaminants by unknown SSOs to hide in the current sample. This issue can partly be resolved by animating each TPF as demonstrated by Kristiansen et al. (in prep.).

## Appendix D: Modified Maximum Likelihood Estimator

As a means to arrive at results efficiently, and as consistency check to the method derived from Wheatland (2004) we fitted  $\alpha$  to the FFDs in  $ED$  and  $E_{Kp}$  space using a Modified Maximum Likelihood Estimator (MMLE, Maschberger & Kroupa 2009). The logarithm of the likelihood function  $\mathcal{L}$  that had to be maximized was given by the authors in Eq. (8) in their manuscript:

$$\log \mathcal{L} = n \log(1 - \hat{\alpha}) - n \log(x_{\max}^{1-\hat{\alpha}} - x_{\min}^{1-\hat{\alpha}}) - \hat{\alpha} \sum_{i=1}^n \log x_i \quad (\text{D.1})$$

where  $x_i$ ,  $x_{\max}$ , and  $x_{\min}$  were the detected flare energies, and the upper and lower limits for detection, respectively.  $n$  was the total number of flares. The estimate for  $\alpha$  would be biased in practice because the value used for  $x_{\max}$  would be the maximum energy that was measured, and not the underlying upper limit. The stabilization transformation suggested by the authors (Eq. (12) in Maschberger & Kroupa 2009) was then applied to the solution for  $\alpha$  to account for this bias:

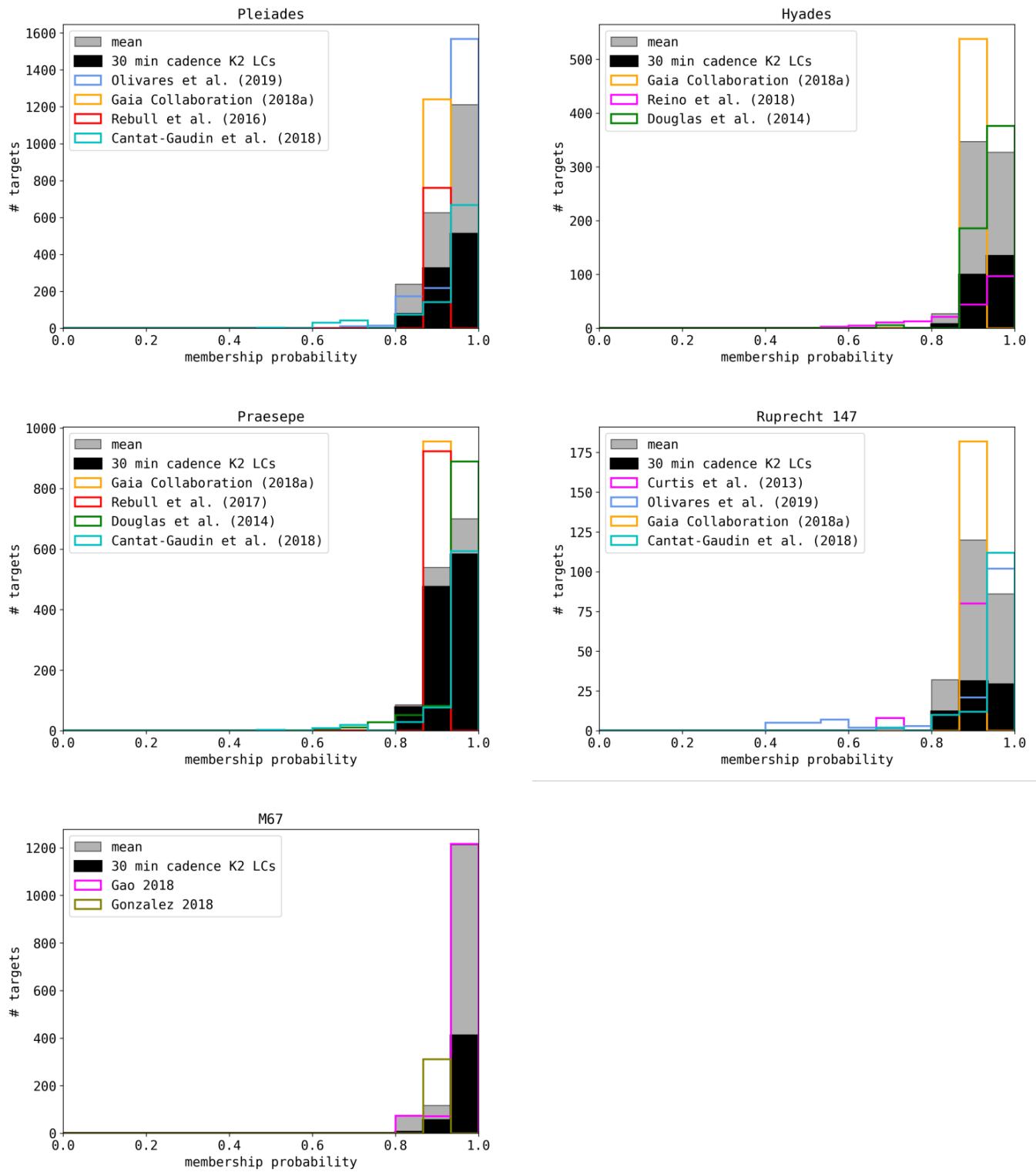
$$\alpha = 1 + \frac{n}{n-2}(\hat{\alpha} - 1) \quad (\text{D.2})$$

Using the MMLE method on the full sample of flares in  $E_{Kp}$  and  $ED$  space we obtained  $\alpha_{\text{erg}} = 1.97$  and  $\alpha_s = 2.06$ , respectively, indicating a marginally flatter power law than the adopted Wheatland (2004) model ( $\alpha_{\text{erg}} = 1.98$ ,  $\alpha_s = 2.1$ ).

**Table A.1.** Membership catalogs used in this work.

source	type	clusters covered	notes
Curtis et al. (2013)	classifier	Ruprecht 147	
Douglas et al. (2014)	probability	Hyades, Praesepe	meta study
Gonzalez (2016)	classifier	M67	
Rebull et al. (2016)	members list	Pleiades	meta study
Rebull et al. (2017)	classifier	Praesepe	meta study
Douglas et al. (2017)	members list	Praesepe	meta study
Gaia Collaboration et al. (2018a)	members list	Hyades <sup>†</sup> , Ruprecht 147, Pleiades, Praesepe	Gaia DR2
Cantat-Gaudin et al. (2018)	probability	Ruprecht 147, Pleiades, Praesepe	Gaia DR2
Gao (2018)	probability	M67	Gaia DR2
Reino et al. (2018)	probability	Hyades <sup>†</sup>	Gaia DR1
Olivares et al. (2018)	probability	Pleiades	Gaia DR2, DANCe
Olivares et al. (2019)	probability	Ruprecht 147	Gaia DR2, DANCe

**Notes.** Meta study: See references in the studies for additional membership information. DANCe: DANCe membership study project. Gaia DR1: The study used data from the first data release (Gaia Collaboration et al. 2016a) of the Gaia mission (Gaia Collaboration et al. 2016b). Gaia DR2: The study used data from the second data release (Gaia Collaboration et al. 2018b). <sup>†</sup>Positions for Hyades were propagated to epoch 2000 using Gaia proper motions.



**Fig. A.1.** Open cluster membership histograms and selected targets. Each panel corresponds to the distribution of candidate members of the cluster indicated in the panel title. Membership catalogs are shown as step histograms (different colors). The mean membership probability distribution is a filled histogram (grey). For our study, we selected only stars with membership probabilities above 80% that were observed by K2 in long cadence mode (black filled histogram).

**Table B.1.** Non-exhaustive literature overview over OC parameters.

cluster	source	distance [pc]	age [Myr]	[Fe/H]
<b>Pleiades</b>	<b>adopted in this work:</b>	135.6	$135 \pm^{25}_{25}$	$-0.037 \pm 0.026$
	Bossini et al. (2019) <sup>a</sup>		$86.5 \pm^{6}_{2.4}$	
	Cantat-Gaudin et al. (2018)	135.6		
	Gossage et al. (2018)		$135 \pm^{25}_{25}$	
	Yen et al. (2018)	126.3	$141.3 \pm^{170}_{100}$	
	Chelli & Duvert (2016)	139		
	Netopil et al. (2016)			-0.01
	Dahm (2015)		$112 \pm^{5}_{5}$	
	Scholz et al. (2015)	130	120	
	Conrad et al. (2014)			$-0.037 \pm 0.026$
	Melis et al. (2014)	136		
	Bell et al. (2012)	135	125	
<b>Hyades</b>	<b>adopted in this work:<sup>c</sup></b>		$690 \pm^{160}_{100}$	$0.13 \pm 0.02$
	Gaia Collaboration et al. (2018a)		$690 \pm^{160}_{100}$	
	Gossage et al. (2018)		680	
	Liu et al. (2016)			$\pm 0.02$
	Netopil et al. (2016)			0.13
	Taylor & Joner (2005)			$0.103 \pm 0.008$
	Salaris et al. (2004)		650	0.15
	Perryman et al. (1998)		$625 \pm^{50}_{50}$	
<b>Praesepe</b>	<b>adopted in this work:</b>	185.5	$750 \pm^{3}_{7}$	0.16
	Bossini et al. (2019)		$750 \pm^{3}_{7}$	
	Cantat-Gaudin et al. (2018)	185.5		
	Gossage et al. (2018)		590	
	Yen et al. (2018)	183	$794 \pm^{253}_{269}$	
	Netopil et al. (2016)			0.16
	Scholz et al. (2015)	187	832	
	Boesgaard et al. (2013)			0.12
	Boudreault et al. (2012)	160	630	
	Salaris et al. (2004)	175	650	
<b>Ruprecht 147</b>	<b>adopted in this work:</b>	305	$2650 \pm^{380}_{380}$	$0.08 \pm 0.07$
	Bragaglia et al. (2018)			$0.08 \pm 0.07$
	Cantat-Gaudin et al. (2018)	305		
	Gaia Collaboration et al. (2018a)	309	$1995 \pm^{404}_{257}$	
	Torres et al. (2018)	283	$2650 \pm^{380}_{380}$	
	Curtis (2016) <sup>b</sup>			$0.10 \pm 0.02$
	Scholz et al. (2015)	270	1953	
	Curtis et al. (2013)	300	$3125 \pm^{125}_{125}$	$0.07 \pm 0.03$
<b>M67</b>	<b>adopted in this work:</b>	908	$3639 \pm^{17}_{17}$	$-0.102 \pm .081$
	Bossini et al. (2019)		$3639 \pm^{17}_{17}$	
	Netopil et al. (2016)			0.03
	Barnes et al. (2016)		$4200 \pm^{700}_{700}$	
	Scholz et al. (2015)		$3428 \pm^{149}_{72}$	
	Conrad et al. (2014)			$-0.102 \pm .081$
	Dias et al. (2012)	908	4300	
	Önehag et al. (2011)	880	4200	0.02
	Salaris et al. (2004)		$4300 \pm^{500}_{500}$	$0.02 \pm 0.06$

**Notes.** <sup>(a)</sup> Bossini et al. (2019) noted some caveats for their determination of ages of young clusters, for which they used Gaia DR2 photometry for isochrone fitting. <sup>(b)</sup> Curtis (2016) reanalysed HIRES spectra using an improved spectroscopic method as compared to Curtis et al. (2013). <sup>(c)</sup> We did not adopt a mean value for the Hyades distance because the cluster members are on average closer than 50 pc.