

HIGH-RESOLUTION MASS MODELS OF DWARF GALAXIES FROM LITTLE THINGS

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ABSTRACT

We present high-resolution rotation curves and mass models of 26 dwarf galaxies from ‘Local Irregulars That Trace Luminosity Extremes, The HI Nearby Galaxy Survey’ (LITTLE THINGS). LITTLE THINGS is a high-resolution ($\sim 6''$ angular; $<2.6 \text{ km s}^{-1}$ velocity resolution) Very Large Array (VLA) HI survey for nearby dwarf galaxies in the local volume within 11 Mpc. The high-resolution HI observations enable us to derive reliable rotation curves of the sample galaxies in a homogeneous and consistent manner. The rotation curves are then combined with *Spitzer* archival 3.6 μm and ancillary optical *U*, *B*, and *V* images to construct mass models of the galaxies. This high quality multi-wavelength dataset significantly reduces observational uncertainties and thus allows us to examine the mass distribution in the galaxies in detail. We decompose the rotation curves in terms of the dynamical contributions by baryons and dark matter halos, and compare the latter with those of dwarf galaxies from The HI Nearby Galaxy Survey (THINGS) as well as Λ CDM Smoothed Particle Hydrodynamic (SPH) simulations in which the effect of baryonic feedback processes is included. Being generally consistent with THINGS and simulated dwarf galaxies, most of the LITTLE THINGS sample galaxies show a linear increase of the rotation curve in their inner regions, which gives shallower logarithmic inner slopes α of their dark matter density profiles. The mean value of the slopes of the 26 LITTLE THINGS dwarf galaxies is $\alpha = -0.32 \pm 0.24$ which is in accordance with the previous results found for low surface brightness galaxies ($\alpha = -0.2 \pm 0.2$) as well as the seven THINGS dwarf galaxies ($\alpha = -0.29 \pm 0.07$). However, this significantly deviates from the cusp-like dark matter distribution predicted by dark-matter-only Λ CDM simulations. Instead our results are more in line with the shallower slopes found in the Λ CDM SPH simulations of dwarf galaxies in which the effect of baryonic feedback processes is included. In addition, we discuss the central dark matter distribution of DDO 210 whose stellar mass is relatively low in our sample to examine the scenario of inefficient supernova feedback in low mass dwarf galaxies predicted from recent Λ CDM SPH simulations of dwarf galaxies where central cusps still remain.

Subject headings: Galaxies: dark matter halos – galaxies: cosmological N–body+SPH simulations – galaxies: kinematics and dynamics

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1. INTRODUCTION

Dark matter the existence of which is indirectly invoked by its gravitational effect in individual galaxies as well as in galaxy clusters, dominates, together with dark energy, the energy budget in the Universe (Zwicky 1937; van den Bergh 1961; Rubin & Ford 1970; Bosma 1978; Peebles 1982; Riess et al. 1998; Perlmutter et al. 1999; Spergel et al. 2003). In particular, the role of dark matter (DM) is critical not only in forming and evolving galaxies in the early Universe but also in shaping the large-scale structure in the Universe through cosmic time (Blumenthal et al. 1984; Colless et al. 2001a; Padmanabhan et al. 2007; Jones et al. 2009; Komatsu et al. 2011 etc.). The cosmological importance of DM has driven efforts to explore the physical nature of DM particles and to attempt their direct detection (Moore et al. 2001; Gaitskell 2004; Angloher et al. 2012; Akerib et al. 2014, and references therein). Of many candidates for DM particles, Cold Dark Matter (CDM) has

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been envisaged as one of the most successful models in that numerical simulations based on a paradigm combining CDM with the cosmological constant, Λ (so-called Λ CDM). This describes well the large-scale structure in the Universe traced by surveys such as the SDSS¹⁷ (York et al. 2000; Doroshkevich et al. 2004), 2dFGRS¹⁸ (Colless et al. 2001b), 6dFGS¹⁹ (Jones et al. 2004), and CMB power spectrum observations (Primack 2003; Spergel et al. 2003, 2007; Planck Collaboration et al. 2014).

However, despite the success of cosmological Λ CDM simulations in producing the large-scale structure of the Universe, distinct differences between the simulations and observations have been found in the DM distribution near the centre of individual galaxies. The simulations have consistently predicted cusp-like DM distributions towards the centers of galaxies, being described by a power law, $\rho \sim R^\alpha$ where R is the galaxy radius and $\alpha \sim -1.0$ (Moore 1994; Navarro et al. 1996b; Navarro et al. 1997; Moore et al. 1999; Ghigna et al. 2000; Klypin et al. 2001; Power et al. 2003; Stoehr et al. 2003; Navarro et al. 2004a; Diemand et al. 2008; Stadel et al. 2009; Navarro et al. 2010; Ishiyama et al. 2013 etc.). In contrast, inferred DM distributions in nearby dwarf galaxies in the local Universe have shown a linear velocity increase towards their centers, giving rise to a sizable density-core ($\rho \sim R^\alpha$ where $\alpha \sim 0.0$) (Moore 1994; de Blok et al. 1996; de Blok & McGaugh 1997; de Blok et al. 2001; de Blok & Bosma 2002; Weldrake et al. 2003; Spekkens et al. 2005; Kuzio de Naray et al. 2006; Kuzio de Naray et al. 2008; Oh et al. 2008; Oh et al. 2011b; Oh et al. 2011a; see, however, Adams et al. 2014 for a discussion on gas kinematics which produces shallower density profiles than those from stellar kinematics). This clear discrepancy of the central DM distribution in galaxies between Λ CDM simulations and observations, the so-called ‘cusp/core’ problem has been one aspect of the small-scale crisis in Λ CDM cosmology which is likely connected to the ‘missing satellites’ problem (Brooks et al. 2013) and ‘too-big-to-fail’ problem (dense satellites; see Boylan-Kolchin et al. 2011, 2012). Given that DM combined with the Λ CDM paradigm is successful in explaining both the large-scale structure in the Universe as well as galaxy formation and evolution, there are good reasons to explore ways to resolve the ‘cusp/core’ problem. We refer to de Blok (2010) and Pontzen & Governato (2014) for the latest review of the ‘cusp/core’ problem.

Late-type dwarf galaxies in the local Universe, with a simple dynamical structure (no bulge and spiral components) have been used for addressing the central DM distribution in galaxies since their dynamics are usually dominated by DM, enabling us to derive more accurately the DM distribution near their centers. These diffuse dwarf galaxies have provided significant observational constraints on the central DM distribution in galaxies. Over the past decade, several high-resolution neutral hydrogen (HI) surveys of galaxies in the local Universe (<11 Mpc) using radio interferometers, among others THINGS²⁰ (Walter et al. 2008), VLA-ANGST²¹

(Ott et al. 2012), LITTLE THINGS²² (Hunter et al. 2012), FIGGS²³ (Begum et al. 2008), SHIELD²⁴ (Cannon et al. 2011), LVHIS²⁵ (Koribalski 2010), have allowed us to derive more reliable HI rotation curves of galaxies and examine their central mass distributions within 1 kpc where the predictions of Λ CDM simulations are most distinctive. For example, high-resolution (0.1–0.2 kpc) DM density profiles of seven dwarf galaxies from THINGS, complemented with the ‘*Spitzer* Infrared Nearby Galaxies Survey’ (SINGS; Kennicutt et al. 2003), were derived by Oh et al. (2008; 2011b). From this, they found that the mean value of the inner density slopes, α , of the seven dwarf galaxies is -0.29 ± 0.07 which is in good agreement with the value of -0.2 ± 0.2 derived earlier from a larger number of Low Surface Brightness (LSB) galaxies (de Blok & Bosma 2002). In the past, it has been argued that observational systematic effects such as beam smearing, center offsets and non-circular motions could have affected the derived central DM distributions of galaxies as these observational biases tend to flatten the derived inner DM density profiles, hiding the central cusps. However, because of the high resolution and quality of the above mentioned observations, these potential biases were significantly reduced. As a result, observational evidence for the core-like distribution of DM near the centers of dIrr galaxies is particularly strong.

In order to explain this behaviour, baryonic feedback processes have been proposed as a means for removing the central cusps expected from Λ CDM DM-only simulations. More specifically, it is expected that DM and baryons in galaxies can be substantially redistributed by frequent explosions of supernovae (SNe) (Larson 1974; Navarro et al. 1996a; Dekel et al. 2003; Mo & Mao 2004; Mashchenko et al. 2008; de Souza et al. 2011; Brook et al. 2011; Pontzen & Governato 2012; Di Cintio et al. 2014; Ogiya & Mori 2014 etc.). However, due to numerical difficulties in simulating multi-phase gas physics as well as the lack of understanding of the detailed baryonic physics in galaxies, taking the baryonic feedback into account in hydrodynamical simulations of dwarf galaxies was considered difficult and limited to simulations of high-redshift galaxies (e.g., Mashchenko et al. 2006).

Governato et al. (2010) were the first to perform high-resolution cosmological N-body+Smoothed Particle Hydrodynamic (SPH) simulations of dwarf galaxies which include the effect of detailed baryonic feedback processes, in particular physically motivated gas outflows driven by SN explosions. From this, they found that the photometric and kinematic properties of the simulated dwarf galaxies are in close, qualitative agreement with those of observed nearby dwarf galaxies. More quantitatively, Oh et al. (2011a) performed an analysis of the baryonic and DM mass distributions of the simulated dwarf galaxies,

[THINGS/Overview.html](#)

²¹ Very Large Array - ACS Nearby Galaxy Survey Treasury (<https://science.nrao.edu/science/surveys/vla-angst>)

²² Local Irregulars That Trace Luminosity Extremes, The HI Nearby Galaxy Survey (<https://science.nrao.edu/science/surveys/littlethings>)

²³ Faint Irregular Galaxies GMRT Survey

²⁴ The Survey of HI in Extremely Low-mass Dwarfs (<http://www.macalester.edu/~jcannon/shield.html>)

²⁵ The Local Volume HI Survey (<http://www.atnf.csiro.au/research/LVHIS/>)

¹⁷ The Sloan Digital Sky Survey (<http://www.sdss.org/>)

¹⁸ The 2dF Galaxy Redshift Survey (<http://www2.aoe.gov.au/2dfgrs/>)

¹⁹ The 6dF Galaxy Survey (<http://oldweb.aoe.gov.au/local/www/6df/>)

²⁰ The HI Nearby Galaxy Survey (<http://www.mpia-hd.mpg.de/>)

in exactly the same way as the THINGS dwarf galaxies were analysed, and showed that their derived rotation curves and the corresponding DM density profiles are consistent with those of the THINGS dwarf galaxies and show a linear increase of velocity in the inner region inherent of shallow DM density profiles. This suggests that repeated gas outflows driven by SN explosions even without a burst of star formation are able to play a fundamental role in removing the central cusps and inducing flatter DM density slopes near the centers of dwarf galaxies. This is in contrast to clusters of galaxies where galaxy interactions are more likely to be the dominant mechanisms for the removal of central cusps rather than star formation activities (Richtler et al. 2011). See Governato et al. (2010) and Pontzen & Governato (2012; 2014) for detailed discussions regarding the effect of SN explosions on the central cusps.

THINGS was only able to probe a small number (i.e., seven) of dwarf galaxies. It is therefore essential to extend the investigation to a larger number of dwarf galaxies in order to obtain a statistically robust observational sample to which simulations can be compared to. Data on a larger sample of dIrr galaxies has now been provided by the latest HI survey of nearby galaxies, LITTLE THINGS (Hunter et al. 2012). LITTLE THINGS is a high-resolution ($\sim 6''$ angular; $\leq 2.6 \text{ km s}^{-1}$ velocity resolution) HI survey for 41 nearby ($< 11 \text{ Mpc}$) gas-rich dwarf galaxies undertaken with the NRAO²⁶ Very Large Array (VLA) in the northern sky. The HI observations are complemented with other wavelength data, such as H α , optical U , B , V , and near infrared (Hunter & Elmegreen 2006), archival *Spitzer* infrared and *GALEX* ultraviolet images as well as follow-up observations with ALMA and *Herschel*. These high-quality multi-wavelength data sets significantly reduce the observational uncertainties inherent in low resolution data which may result in hiding the central cusps. Of the 41 galaxies, we select a sample of 26 dwarf galaxies (three of them are also in THINGS) which show a regular rotation pattern in their velocity fields. In this paper, we extract (1) bulk and non-circular motions of the sample galaxies, (2) derive rotation curves, (3) decompose the derived rotation curves in terms of the contributions by baryons and DM halos, and (4) address the central DM distribution by making a direct comparison between the derived DM distributions of the galaxies to those of SPH+N-body simulations of dwarf galaxies.

The structure of this paper is as follows. The data used for deriving the mass models of our sample galaxies are described in Section 2. In Section 3, we present the rotation curves, the mass models of baryons, and the DM mass modeling of the galaxies. In Section 4, we discuss the central DM distributions of the sample galaxies by comparing them with those of dwarf galaxies from both THINGS and simulations. We then discuss the effect of SN feedback on the central cusp in Section 5, followed by the discussion of the effect of beam smearing on the cen-

²⁶ NRAO is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. These data were taken during the upgrade of the VLA to the Expanded VLA or EVLA. In this paper we refer to the instrument as the VLA, the retrofitted antennas as EVLA antennas, and non-retrofitted antennas as VLA antennas. This emphasizes the hybrid nature of the instrument and distinguishes it from the far more powerful Jansky VLA or JVLA it has become since 2012.

tral DM distribution of galaxies in Section 6. Lastly, we summarise the main results of this paper and conclusions in Section 7.

2. THE DATA

We use high-resolution HI data of 26 nearby ($< 11 \text{ Mpc}$) dwarf galaxies from LITTLE THINGS to address the central DM distribution of the galaxies. The sample galaxies show a regular rotation pattern in their 2-dimensional (2D) HI velocity fields (see Appendix), allowing us to derive reliable rotation curves which include the contributions to their kinematics of both their DM halo, i.e., non-baryonic and baryonic matter. Considering the distances of the sample galaxies, the linear resolutions of the LITTLE THINGS HI data ($\sim 6''$) range from ~ 26 to 200 pc with an average of 100 pc which is sufficient to resolve the inner 1 kpc region of the galaxies in so far unmatched detail. This enables us to examine the central DM distributions of the galaxies in detail. In addition, observational systematic effects inherent in low-resolution data (e.g., beam smearing, kinematic center offset and non-circular motions) are significantly reduced in the high-resolution HI data, which allows us to derive more accurate underlying kinematics of the sample galaxies.

Although the total kinematics of late-type dwarf galaxies is dominated by DM (Prada & Burkert 2002), it is nonetheless important to separate the contribution by baryons from the total rotation curve. This is achieved by using *Spitzer* archival *IRAC* $3.6\mu\text{m}$ and ancillary optical color information (Hunter & Elmegreen 2006). *Spitzer* *IRAC* $3.6\mu\text{m}$ images are much less affected by dust than maps at shorter wavelengths and trace the old stellar populations that occupy the dominant fraction of the stellar mass in galaxies (Walter et al. 2007). Although there is some contamination by PAH emission extending into the $3.6\mu\text{m}$ band (Meidt et al. 2014) this is reduced in dIrr galaxies which, by virtue of having low heavy element abundances, have a correspondingly lower dust content.

According to the simulations by Governato et al. (2012), the degree of baryonic feedback is largely dependent on the amount of stars, so past star formation activity in the sense that for the same total mass budget (baryons + DM), those galaxies with a higher past star formation activity will have had more significant outflows. Therefore a reliable measurement of the stellar mass in a galaxy is essential to investigating the effect of baryonic feedback on the central cusp. The basic observational properties of the sample galaxies are listed in Table 1. We refer to Hunter et al. (2012) for a complete description of the HI observations and data reduction.

3. MASS MODELS

In this Section, we perform mass modeling of the 26 LITTLE THINGS dwarf galaxies using the high-resolution VLA HI data and *Spitzer* *IRAC* $3.6\mu\text{m}$ images as described in Section 2. As mentioned earlier, the simple dynamical structure and dominant circular rotation in the disk of the sample galaxies help to reduce uncertainties stemming from the decomposition of galaxy rotation curves into contributions due to baryonic and non-baryonic matter.

The first step in the mass modelling includes deriv-

TABLE 1
PROPERTIES AND TILTED-RING PARAMETERS OF THE LITTLE THINGS SAMPLE GALAXIES

Name	α (2000.0)	δ (2000.0)	D (Mpc)	V_{sys} (km s $^{-1}$)	(P.A.) ($^{\circ}$)	(i) ($^{\circ}$)	M_V	12 + log(O/H)	log(SFR $_{D}^{H\alpha}$) (M $_{\odot}$ yr $^{-1}$ kpc $^{-2}$)	log(SFR $_{D}^{FUV}$) (M $_{\odot}$ yr $^{-1}$ kpc $^{-2}$)
	(h m s)	($^{\circ}$, $'$, $''$)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
CVnIdwA	12 38 39.2	+32 45 41.0	3.6	306.2 \pm 1.3	48.4 \pm 13.9	66.5 \pm 5.2	-12.4	7.3 \pm 0.06	-2.58 \pm 0.01	-2.48 \pm 0.01
DDO 43	07 28 17.7	+40 46 08.3	7.8	355.4 \pm 3.6	294.1 \pm 0.1	40.6 \pm 0.1	-15.1	8.3 \pm 0.09	-1.78 \pm 0.01	-1.55 \pm 0.01
DDO 46	07 41 26.3	+40 06 37.5	6.1	360.8 \pm 1.3	274.1 \pm 5.0	27.9 \pm 0.1	-14.7	8.1 \pm 0.10	-2.89 \pm 0.01	-2.46 \pm 0.01
DDO 47	07 41 55.3	+16 48 07.1	5.2	270.7 \pm 1.3	311.6 \pm 11.9	45.5 \pm 9.0	-15.5	7.8 \pm 0.20	-2.70 \pm 0.01	-2.40 \pm 0.01
DDO 50	08 19 03.7	+70 43 24.6	3.4	156.5 \pm 1.2	175.7 \pm 10.1	49.7 \pm 6.0	-15.6	7.7 \pm 0.14	-1.67 \pm 0.01	-1.55 \pm 0.01
DDO 52	08 28 28.4	+41 51 26.5	10.3	399.0 \pm 2.3	8.2 \pm 5.3	43.0 \pm 0.0	-15.4	7.7 \pm ...	-3.20 \pm 0.01	-2.43 \pm 0.01
DDO 53	08 34 06.4	+66 10 47.9	3.6	18.6 \pm 0.7	131.6 \pm 0.1	27.0 \pm 0.0	-13.8	7.6 \pm 0.11	-2.42 \pm 0.01	-2.41 \pm 0.01
DDO 70	10 00 00.9	+05 20 12.9	1.3	303.5 \pm 1.2	44.5 \pm 13.5	50.0 \pm 0.0	-14.1	7.5 \pm 0.06	-2.85 \pm 0.01	-2.16 \pm 0.00
DDO 87	10 49 34.9	+65 31 47.9	7.7	340.0 \pm 1.0	235.1 \pm 3.9	55.5 \pm 4.8	-15.0	7.8 \pm 0.04	-1.36 \pm 0.01	-1.00 \pm 0.01
DDO 101	11 55 39.1	+31 31 9.9	6.4	589.4 \pm 1.0	287.4 \pm 2.8	51.0 \pm 7.1	-15.0	8.7 \pm 0.03	-2.85 \pm 0.01	-2.81 \pm 0.01
DDO 126	12 27 06.6	+37 08 15.9	4.9	219.4 \pm 1.8	138.0 \pm 3.6	65.0 \pm 0.0	-14.9	7.8 \pm ...	-2.37 \pm 0.01	-2.10 \pm 0.01
DDO 133	12 32 55.2	+31 32 19.1	3.5	330.7 \pm 1.0	359.6 \pm 8.4	43.4 \pm 0.1	-14.8	8.2 \pm 0.09	-2.88 \pm 0.01	-2.62 \pm 0.01
DDO 154	12 54 05.7	+27 09 09.9	3.7	372.0 \pm 1.3	226.3 \pm 3.1	68.2 \pm 3.1	-14.2	7.5 \pm 0.09	-2.50 \pm 0.01	-1.93 \pm 0.01
DDO 168	13 14 27.3	+45 55 37.3	4.3	192.6 \pm 1.2	275.5 \pm 5.8	46.5 \pm 0.1	-15.7	8.3 \pm 0.07	-2.27 \pm 0.01	-2.04 \pm 0.01
DDO 210	20 46 51.6	-12 50 57.7	0.9	-139.5 \pm 1.0	65.0 \pm 0.0	66.7 \pm 0.1	-10.9	7.2 \pm	-2.71 \pm 0.06
DDO 216	23 28 34.7	+14 44 56.2	1.1	-186.0 \pm 1.1	133.6 \pm 7.5	63.7 \pm 4.6	-13.7	7.9 \pm 0.15	-4.10 \pm 0.07	-3.21 \pm 0.01
F564-V3	09 02 54.0	+20 04 26.3	8.7	481.0 \pm 2.0	122.5 \pm 0.0	56.5 \pm 9.9	-14.0	7.6 \pm	-2.79 \pm 0.02
IC 10	00 20 18.9	+59 17 49.9	0.7	-348.0 \pm 2.9	55.7 \pm 10.1	47.0 \pm 15.6	-16.3	8.2 \pm 0.12	-1.11 \pm 0.01	...
IC 1613	01 04 49.6	+02 48 14.1	0.7	-232.2 \pm 2.2	48.0 \pm 0.0	7.6 \pm 0.05	-14.6	7.6 \pm 0.05	-2.56 \pm 0.01	-1.99 \pm 0.01
NGC 1569	04 30 46.2	+64 51 10.3	3.4	-85.4 \pm 5.6	122.5 \pm 1.5	69.1 \pm 0.1	-18.2	8.2 \pm 0.05	+0.19 \pm 0.01	-0.01 \pm 0.01
NGC 2366	07 28 53.4	+69 12 49.6	3.4	103.5 \pm 1.4	38.7 \pm 4.2	63.0 \pm 0.8	-15.8	7.9 \pm 0.01	-1.67 \pm 0.01	-1.66 \pm 0.01
NGC 3738	11 35 46.9	+54 31 44.8	4.9	235.8 \pm 1.2	292.2 \pm 5.1	22.6 \pm 0.1	-17.1	8.4 \pm 0.01	-1.66 \pm 0.01	-1.53 \pm 0.01
UGC 8508	13 30 44.9	+54 54 32.4	2.6	60.5 \pm 1.1	126.1 \pm 3.2	82.5 \pm 0.1	-13.6	7.9 \pm 0.20	-2.03 \pm 0.01	...
WLM	00 01 59.9	-15 27 57.2	1.0	-122.3 \pm 1.2	174.5 \pm 2.8	74.0 \pm 0.1	-14.4	7.8 \pm 0.06	-2.77 \pm 0.01	-2.05 \pm 0.01
Haro 29	12 26 18.4	+48 29 40.4	5.9	279.3 \pm 2.2	214.5 \pm 2.8	61.2 \pm 4.0	-14.6	7.9 \pm 0.07	-0.77 \pm 0.01	-1.07 \pm 0.01
Haro 36	12 46 56.6	+51 36 47.3	9.3	499.3 \pm 3.9	248.4 \pm 12.5	70.0 \pm 0.0	-15.9	8.4 \pm 0.08	-1.86 \pm 0.01	-1.55 \pm 0.01

Note. (1)(2): Kinematic center position derived from the tilted-ring analysis in Section 3.1.1; (3): Distance as given in Hunter et al. (2012); (4): Systemic velocity derived from the tilted-ring analysis in Section 3.1.1; (5): Average value of the kinematic position angle (PA) derived from the tilted-ring analysis in Section 3.1.1. PA is the angle measured counter-clockwise from the north direction in the sky to the major axis of the receding half of the galaxy; (6): Average value of the kinematic inclination derived from the tilted-ring analysis in Section 3.1.1; (7): Absolute V magnitude as given in Hunter et al. (2012); (8): Oxygen abundance taken from the literature as compiled in Hunter et al. (2012). ‘...’ indicates no uncertainty available on the measurements; (9)(10): The $H\alpha$ and $GALEx FUV$ star formation rates normalized to the area (πR_d^2 where R_d is the disk scale length) as given in Hunter & Elmegreen (2004) and Hunter et al. (2010), respectively. ‘...’ indicates that no measurements are available.

TABLE 2
MASS MODELLING RESULTS OF THE LITTLE THINGS SAMPLE GALAXIES

Name	R_{\max} (kpc)	$R_{0,3}$ (kpc)	V_{\max} (km s ⁻¹)	$V_{\text{ISO}}(R_{\max})$ (km s ⁻¹)	$R_{\max} \text{H}_\text{I-beam}^{-1}$ (kpc)	z_0	c	V_{200} (km s ⁻¹)	R_c (kpc)	$(10^{-3} M_\odot \text{ pc}^{-3})^{\alpha_0}$	α_{\min}	$\alpha_{\text{TF},pdf}$	M_{gas} ($10^7 M_\odot$)	M_{IR} ($10^7 M_\odot$)	M_{\star} ($10^7 M_\odot$)	$\log(M_{200})$ (M_\odot)	
CVnIdwA	2.59	2.27	23.5	23.7	13.5	0.24	-0.4 (11.0)	478.5 (16.1 ± 1.6)	2.01 ± 0.52	8.19 ± 1.62	-1.25 ± 0.21	+0.03 ± 0.27	2.91	0.41	0.49	8.529	9.138
DDO 43	4.19	4.19	33.2	38.5	13.7	0.24	5.6 ± 2.6 (10.2)	39.8 ± 14.9 (31.4 ± 1.8)	0.94 ± 0.18	33.22 ± 8.78	-0.15 ± 0.13	-0.25* ± 0.16	23.26	9.159	10.317
DDO 46	2.92	2.04	66.1	77.4	15.7	...	18.7 ± 3.5 (9.0)	61.2 ± 9.7 (136.7 ± 6.1)	0.51 ± 0.04	517.82 ± 63.46	-0.41 ± 0.01	-0.42* ± 0.02	22.08	9.609	11.925
DDO 47	7.71	1.55	64.7	68.9	29.3	...	-0.4 (9.2)	1359.0 (51.3 ± 2.5)	3.72 ± 0.48	12.15 ± 1.49	-1.24* ± 0.20	-1.25* ± 0.22	46.80	9.930	10.648
DDO 50	9.81	1.67	35.7	84.4	0.28	22.2 (10.3)	15.4 (28.8 ± 0.2)	0.15 ± 0.08	379.64 ± 381.14	-0.41 ± 0.41	+0.10 ± 0.41	132.52	9.463	9.895	
DDO 52	5.43	60.2	60.5	15.9	6.1 ± 1.1 (9.4)	71.9 ± 11.8 (57.8 ± 1.3)	1.33 ± 0.07	48.81 ± 3.63	-0.55 ± 0.04	-0.49 ± 0.02	33.43	7.20	10.60	5.31	10.803		
DDO 53	1.45	0.62	28.6	33.1	13.1	0.14	0.0 ± 4.4 (10.4)	425.8 ± 2531.9 (27.8 ± 3.2)	2.22 ± 1.95	25.10 ± 6.63	-0.27 ± 0.48	+0.14 ± 0.80	7.00	0.96	0.97	8.567	9.849
DDO 70	2.00	0.55	43.9	35.5	22.9	0.11	0.0 ± 10.2 (10.3)	866.1 ± 11717.4 (38.6 ± 1.8)	0.51 ± 0.11	119.95 ± 34.69	-0.43 ± 0.02	-0.48 ± 0.03	3.80	1.24	1.96	8.768	10.277
DDO 87	7.39	4.13	56.6	26.1	0.48	0.0 ± 1.2 (10.2)	511.4 ± 143.8 (40.7 ± 1.6)	0.46 ± 0.11	139.0 ± 9.77	-0.01 ± 0.44	-0.01 ± 0.47	29.12	6.18	3.27	9.734	10.346	
DDO 101	1.95	1.05	64.5	63.6	7.5	0.13	25.0 ± 1.3 (9.4)	43.9 ± 115.7 (6.4)	0.32 ± 0.01	849.14 ± 77.23	-1.00 ± 0.15	-1.02 ± 0.12	3.48	5.79	6.54	9.279	51.018
DDO 126	3.99	2.88	37.2	40.5	24.3	0.23	-0.2 ± 44.4 (10.1)	626.6 ± 48442.1 (31.4 ± 1.2)	1.33 ± 0.10	21.55 ± 2.00	-0.39 ± 0.15	-0.39 ± 0.15	16.36	2.27	1.62	9.182	10.008
DDO 135	3.48	3.48	46.1	47.5	16.5	0.22	4.9 ± 23.9 (9.8)	74.5 ± 36.7 (46.6 ± 1.6)	0.83 ± 0.06	73.69 ± 7.61	-0.11 ± 0.15	-0.11 ± 0.15	12.85	2.62	3.04	9.261	10.522
DDO 154	7.32	2.59	47.8	49.2	51.3	0.14	6.4 ± 15 (9.8)	38.7 ± 2.9 (41.1 ± 0.6)	0.95 ± 0.03	53.21 ± 3.19	-0.39 ± 0.11	-0.41 ± 0.13	35.27	1.31	0.83	9.614	10.359
DDO 168	3.14	60.3	67.4	19.2	0.22	-0.3 ± 117.6 (9.2)	229.0 ± 8.4 (63.9 ± 4.0)	2.81 ± 0.83	48.25 ± 2.45	-0.28 ± 0.28	-0.97 ± 0.18	25.94	5.13	5.85	9.520	10.934	
DDO 210	0.31	12.0	11.3	6.0	0.10	0.0 ± 292.3 (12.2)	339.4 ± 10 ... (12.4 ± 1.0)	0.29 ± 0.05	116.43 ± 23.35	-0.30 ± 0.07	-0.70 ± 0.02	0.14	0.04	0.06	6.964	8.797	
DDO 216	1.12	0.48	18.9	15.7	13.0	0.27	13.2 ± 2.7 (11.7)	8.8 ± 1.0 (14.7 ± 0.4)	0.15 ± 0.03	127.02 ± 43.35	-0.17 ± 0.47	-0.30 ± 0.62	0.49	1.51	1.51	7.807	9.019
F564-V3	3.71	1.83	34.2	36.5	27.0	0.07	10.8 ± 4.5 (10.4)	26.2 ± 7.1 (29.6 ± 2.1)	0.55 ± 0.20	74.24 ± 40.17	-0.66* ± 0.14	-0.68* ± 0.18	4.37	9.003	9.931
IC 10	0.54	0.20	36.5	36.1	27.0	0.07	0.2 ± 95.1 (18.6)	2077.1 ± ... (36.8 ± 2.2)	0.27 ± 0.12	190.4 ± 76.74	-1.19 ± 0.01	-0.25 ± 0.32	1.65	11.81	...	8.213	10.215
IC 1613	2.71	0.36	16.9	21.1	103.8	0.15	5.0 ± ... (11.2)	5.4 ± ... (12.5 ± 0.4)	0.20 ± 0.04	19.25 ± 3.45	+1.16 ± 1.22	0.10 ± 0.32	5.93	1.94	2.88	8.448	8.808
NGC 1569	3.05	29.1	46.9	31.5	0.19	-0.3 ± 17.2 (9.8)	736.6 ± ... (52.6 ± 2.2)	2.71 ± 0.81	15.25 ± 2.45	-0.77 ± 0.17	-0.23 ± 0.67	20.24	20.69	36.29	9.193	10.352	
NGC 2366	8.08	2.75	58.2	60.8	70.7	0.29	5.2 ± 0.6 (9.4)	66.0 ± 6.1 (52.6 ± 0.8)	1.21 ± 0.04	43.89 ± 2.51	-0.52 ± 0.13	-0.53 ± 0.12	108.24	10.81	6.94	9.841	10.680
NGC 3738	1.75	1.10	125.5	130.2	11.7	0.12	12.4 ± 15.4 (8.1)	310.4 ± 536.2 (598.2 ± 30.4)	0.45 ± 0.04	2132.36 ± 277.75	-0.44 ± 0.02	-0.42 ± 0.03	12.58	12.48	46.62	9.838	13.848
UGC 8508	1.86	0.46	46.0	43.5	25.0	0.09	7.8 ± 1.2 (10.0)	54.7 ± 7.4 (44.2 ± 2.2)	1.96 ± 0.21	45.29 ± 2.38	-0.38 ± 0.16	-0.30 ± 0.01	1.19	0.30	0.77	8.913	10.453
WLM	3.04	1.60	37.3	37.7	82.9	3.0 ± 1.0 (8.2)	90.8 ± 26.1 (33.3 ± 0.7)	1.74 ± 0.01	57.46 ± 1.57	+0.03 ± 0.01	-0.02 ± 0.01	7.96	1.23	9.002	10.085	10.085	
Haro 29	5.03	0.64	30.5	35.0	25.7	0.13	19.5 ± 1.9 (10.4)	23.4 ± 0.9 (32.7 ± 0.9)	0.23 ± 0.04	414.68 ± 137.20	-0.50 ± 0.12	-0.50* ± 0.12	9.35	1.43	9.156	10.061	10.061
Haro 36	3.16	58.2	52.6	10.0	0.13	-0.4 ± 1503.4 (9.7)	1649.2 ± ... (47.9 ± 3.5)	8.40 ± 12.17	16.99 ± 2.78	-0.48 ± 0.06	-0.50 ± 0.01	11.16	5.81	...	9.308	10.558	

(1): The radius where the outermost part of the rotation curve is measured; **(2):** The radius where the logarithmic slope of the total rotation curve (DM + baryons) $d\log V/d\log R = 0.3$. See Section 4.1; **(3):** The radius where the logarithmic slope of the best fitted pseudo-isothermal halo model at R_{\max} ; **(4):** The rotation velocity (asymmetric drift corrected) at R_{\max} ; **(5):** The ratio of R_{\max} to H_I beam size; **(6):** The vertical scale height of the stellar disk; **(7)(8):** Concentration parameter c and rotation velocity V_{200} for a NFW halo model (Navarro et al. 1995; 1996b); **c** values in brackets are derived using an empirical relationship between $c - V_{200}$ as given by McGaugh et al. (2007). The corresponding V_{200} values are fitted after fixing c to the ones in brackets assuming a ‘minimum disk’. See Section 3.4 for a detailed description. ‘...’ indicates that the uncertainty is unphysically large; **(9)(10):** Core-radius and core-density of a pseudo-isothermal halo model (Begeman et al. 1991); **(11):** The logarithmic inner slope α_{\min} of the total matter (DM halo + baryons) density profiles measured in Section 4.2; **(12):** The logarithmic inner slope $\alpha_{\text{TF},pdf}$ of the DM density profiles measured in Section 4.2; **(13):** The gas mass derived in Section 3.2; **(14):** The stellar mass derived from the kinematic analysis in Section 3.3; **(15):** The stellar mass derived using a spectral energy distribution (SED) fitting technique in Zhang et al. (2012); **(16):** The dynamical mass measured using $V_{\text{ISO}}(R_{\max})$ and R_{\max} ; **(17):** CDM halo mass M_{200} estimated from the V_{200} in brackets of column (5) and using Eq. 3 given in Oh et al. (2011a). See Section 3.4 for more details.

*The dynamical contribution by the stellar component is included in the DM density profile as no *Spitzer* 3.6μm image is available for this galaxy.

ing rotation curves of the sample galaxies using the 2D HI velocity fields which reflect the total kinematics including both baryons and a DM halo. HI is mostly distributed in the disk of a galaxy where the circular rotation is dominant and is a useful kinematic tracer for deriving the galaxy rotation curve. This is mainly due to the larger radial extent of HI in the disk compared to stellar components (e.g., 3–4 times; Sofue & Rubin 2001). We proceed to derive mass models of the baryons (gas and stars) using HI integrated intensity maps and *Spitzer IRAC* 3.6 μm images, and subtract their contribution from the total kinematics. Lastly, we quantify the kinematic residuals in order to examine the DM distribution near the centers of the galaxies. In the following sections, we describe these mass modelling procedures in more detail.

3.1. Rotation curves: total kinematics

3.1.1. Tilted-ring fits

For the derivation of the rotation curves of our sample galaxies, we fit a 2D tilted-ring model which consists of a series of concentric ellipses to the 2D velocity fields extracted from the HI data cubes of the sample galaxies (Rogstad et al. 1974). Each ellipse has its own geometric and kinematic parameters, such as center position (XPOS, YPOS), position angle (PA), inclination (INCL), systemic velocity (VSYS), and rotation velocity (VROT). This so-called ‘tilted-ring analysis’ of 2D velocity fields obtained from HI, CO or H α spectroscopic observations has been widely used for deriving rotation curves of disk-dominated galaxies (e.g., rotation curves of the THINGS galaxies sample; de Blok et al. 2008).

Tilted-ring models, however, only hold for those cases where the velocity field is a reliable representation of the overall kinematics of a galaxy. Non-circular motions caused by star-forming activity like stellar winds, SNe, etc., as well as spiral arms, a bar-like or triaxial galaxy potential, galaxy mergers or tidal interactions disturb gas motions in galaxies on small and large scales, resulting in distorted velocity fields.

Low resolution HI data usually smooth any features related to small-scale non-circular motions. This is known as beam smearing and tends to yield a velocity gradient along the major axis of a galaxy that is less steep, particularly in the central regions (e.g., de Blok & McGaugh 1997; Swaters et al. 2000; van den Bosch & Swaters 2001; McGaugh et al. 2001). Beam smearing is significantly reduced in high-resolution data such as THINGS and LITTLE THINGS (see de Blok et al. 2001; McGaugh et al. 2001; Kuzio de Naray et al. 2006).

The LITTLE THINGS sample dwarf galaxies selected in this study, with few exceptions, appear to show no significant large-scale kinematic features, such as bars, spiral arms or warps in their HI or *Spitzer IRAC* 3.6 μm images. However, our sample galaxies do suffer from the effect of small-scale turbulent gas motions due to stellar winds and SNe. In general, the derived velocity field in dwarf galaxies is more vulnerable to the impact of stellar activity due to their low gravitational potential (Walter et al. 1998; Walter & Brinks 2001; see also Bagetakos et al. 2011). Therefore, the extraction of a robust velocity field is essential if one wants to derive the undisturbed underlying kinematics of a galaxy.

In order to correct for small-scale random motions and extract only the component due to circular rotating velocity in a galaxy, we derive the ‘bulk velocity field’ as proposed by Oh et al. (2008; 2011b). Compared with other typical types of velocity fields, such as intensity-weighted mean (IWM), single Gaussian fit and hermite h3, the bulk velocity field has been found to be ideally suited to extracting the underlying bulk rotation of a galaxy in the presence of random non-circular motions. We point out, as an example, how the ‘kinks’ or ‘wiggles’ of the iso-velocity contours of the extracted bulk velocity fields displayed in panel (e) of Fig. A.1 are weaker than those of the IWM velocity fields (panel b). The bulk velocity field appears to much better represent the overall kinematics of the galaxy and hence the underlying gravitational potential.

Following the standard procedure described in Begeaman (1989), we fit tilted-ring models to the bulk velocity fields of the sample galaxies. For this, we use the ‘rotcur’ task in GIPSY²⁷ (van der Hulst et al. 1992). The derived rotation curves of the galaxies are presented in the figures of the Appendix (e.g., A.2). As seen from the scatter in the fits made with all ring parameters free (open circles) in the figures, the extracted bulk velocity fields are not completely free from the effect of small-scale random motions in the galaxies. However, they are relatively insignificant and are averaged out after several iterations as shown in the final rotation curves (solid lines) in the figures.

This is also confirmed in the harmonic analysis of the velocity fields. As described in Schoenmakers et al. (1997), we perform harmonic decompositions of the bulk and IWM velocity fields of the sample galaxies. For this, we use the task ‘RESWR’ in GIPSY. We expand the velocity fields into *sine* and *cosine* terms up to 3rd order (i.e., c_m and s_m where $m = 1, 2$ and 3) after fixing the center position, PA and INCL with those derived from the tilted-ring analysis. If we allow RESWR to fit a velocity field with center position, PA and INCL as free parameters, non-circular motions tend to be absorbed into variations in these geometrical parameters, underestimating the amount of non-circular motions. As an example, streaming non-circular motions in a barred galaxy are mainly responsible for the radial motions that are typically reflected in the s_1 and s_3 terms (Schoenmakers et al. 1997; Wong et al. 2004; Spekkens & Sellwood 2007). However, these radial motions can also be modelled by a radial variation of PA, without the need for s_1 and s_3 terms. Similarly, other ring parameters can affect the harmonic analysis in the same way if they are kept as free parameters in the fit. To quantify and describe non-circular motions, we calculate the absolute amplitudes $\langle A \rangle$ and the phases of each component decomposed. For the amplitudes, we take the median of $A_m(R)$ as described in Schoenmakers et al. (1997) (see also Trachternach et al. 2008),

for $m = 1$,

$$A_1(R) = \sqrt{s_1(R)^2}, \quad (1)$$

for $m > 1$,

$$A_m(R) = \sqrt{c_m(R)^2 + s_m(R)^2}, \quad (2)$$

²⁷ The Groningen Image Processing System

where R is the galaxy radius.

As shown in the section labelled ‘*Harmonic Analysis*’ in the Appendix, the amplitudes of harmonic terms (e.g., c_2 , s_1 and s_2 which are corrected for inclination) derived from the bulk velocity fields (black dots in the Harmonic Analysis panels) are lower than those derived using the IWM velocity fields over all radii. This shows that the effect of random non-circular motions are largely reduced in the bulk velocity fields. The tilted-ring parameters of the sample galaxies derived using the bulk velocity fields are given in Table 1. We note that the kinematic center positions given in Table 1 that are derived from the tilted-ring analysis are offset from the optical (V -band, mostly) central isophot used by Hunter et al. (2012).

3.1.2. Asymmetric drift correction

Pressure support caused by random gas motions in the gaseous disk of a galaxy tends to lower the rotation velocity, which results in an underestimate of the dynamical mass of the galaxy. In general the dynamical effect of the pressure support is higher in the outer region of a galaxy where the gas density is low. In particular, this, so-called ‘asymmetric drift’ is significant in dwarf galaxies whose maximum rotation velocities are comparable to the velocity dispersions in the gas disk. The asymmetric drift correction hence should be made to derive more reliable rotation curves of galaxies where the dynamical support by random motions to its gas disk is significant. This is the case of some of our sample galaxies whose velocity dispersions are comparable to their maximum rotation velocities in the outer region. For the asymmetric drift correction, we follow the method described in Bureau & Carignan (2002; see also Oh et al. 2011b) as follows:

The asymmetric drift correction σ_D is given as,

$$\begin{aligned} \sigma_D^2 &= -R\sigma^2 \frac{\partial \ln(\rho\sigma^2)}{\partial R} \\ &= -R\sigma^2 \frac{\partial \ln(\Sigma_{\text{HI}}\sigma^2)}{\partial R}, \end{aligned} \quad (3)$$

where σ_D is the asymmetric drift correction, R is the galaxy radius, σ is the velocity dispersion, and ρ is the volume density of gas disk. In general, ρ can be approximated as the gas surface density Σ_{HI} for a gas disk with an exponential distribution in the vertical distribution and a constant scale height z_0 (i.e., $d(\ln(z_0))/dr = 0$). Large fluctuations in the derivative in Eq. 3 can be smoothed by fitting an analytical function which has three free parameters, I_0 [$M_\odot \text{ pc}^{-2} \text{ km}^2 \text{ s}^{-2}$], R_0 [arcsec], and α [arcsec^{-1}] to the numerator as follows,

$$\Sigma_{\text{HI}}\sigma^2(R) = \frac{I_0(R_0 + 1)}{R_0 + e^{\alpha R}}. \quad (4)$$

Lastly, the corrected rotation velocity V_{cor} is derived by adding the asymmetric drift correction σ_D to the rotation velocity V_{rot} derived from tilted-ring fits, quadratically,

$$V_{\text{cor}}^2 = V_{\text{rot}}^2 + \sigma_D^2. \quad (5)$$

The analytical function given in Eq 4 provides a good fit to most sample galaxies except for DDO 52, IC 10, NGC 3738 and UGC 8508 where a significant degree of radial fluctuation is present in $\Sigma_{\text{HI}}\sigma^2$ at small radii. However, the correction at small radii is insignificant and thus will not affect significantly the final results.

Valenzuela et al. (2007) have shown that the standard asymmetric correction can be underestimated without considering the gas pressure gradients triggered by star formation and feedback. However, the effect of additional thermal pressure gradients in the gas is most likely insignificant for the sample galaxies in this study. The thermal pressure is already included in the observed velocity dispersion used for the asymmetric drift correction as part of its broadening. In addition, as shown in the velocity dispersion map (i.e., moment 2) in the Appendix, significant anisotropy in the dispersion is not found in the galaxies. This implies that any separation of turbulent and thermal pressures, with explicit calculation of asymmetries in the turbulent component, would have insignificant effect on the asymmetric drift corrected rotation curves. The asymmetric drift corrected rotation curves of the sample galaxies which are used for the mass modelling are presented in the Appendix.

3.2. Gas distribution

The rotation curves derived in Section 3.1 already provide a good approximation of the DM halos’ kinematics of the sample galaxies given the dominant contribution of the DM halo to the total kinematics of dwarf galaxies. However, to derive more accurate DM distributions of the galaxies, we construct mass models of their gaseous and stellar components which account for most of the baryons in dwarf galaxies.

We use total integrated HI intensity maps (moment 0) of the galaxies to derive the mass model of the gaseous component. For consistency with the rotation curves in Section 3.1, we apply the derived tilted-ring models to the HI intensity maps, and obtain gas surface density profiles of the galaxies which are scaled up by a factor of 1.4 to take Helium and metals into account. We then convert the gas surface density profiles to the corresponding gas rotation velocities assuming that gas components are mainly distributed in a thin disk. As an example, the derived gas surface density profile and the corresponding rotation velocity of CVnIdwA are shown in the panels (g) and (h) of Fig. A.3, respectively. Here, we do not correct for the effect of molecular hydrogen (H_2) since low metallicities in dwarf galaxies can induce only a small fraction of the gaseous component in the form of H_2 (e.g., Leroy et al. 2007; Leroy et al. 2011; Schruba et al. 2012).

3.3. Stellar distribution

We use *Spitzer IRAC* 3.6 μm images to derive mass models of the stellar components of the galaxies. Compared with optical images, the *Spitzer IRAC* 3.6 μm image is less affected by dust and less sensitive to young stellar populations which usually emit most energy in the optical regime but occupy only a small fraction of the total stellar mass. Instead, the *Spitzer IRAC* 3.6 μm image is useful for tracing old stellar populations that are dominant in late-type dwarf galaxies. This enables us to derive a robust estimate of the stellar mass of our sample galaxies as used for deriving the mass models of the stellar components of THINGS galaxies (de Blok et al. 2008; Trachternach et al. 2008; Oh et al. 2008, 2011b).

Like we did for the gas component, we derive 3.6 μm surface brightness profiles of the stellar components of the galaxies by applying the derived tilted-ring parameters to the *Spitzer IRAC* 3.6 μm images as shown in the

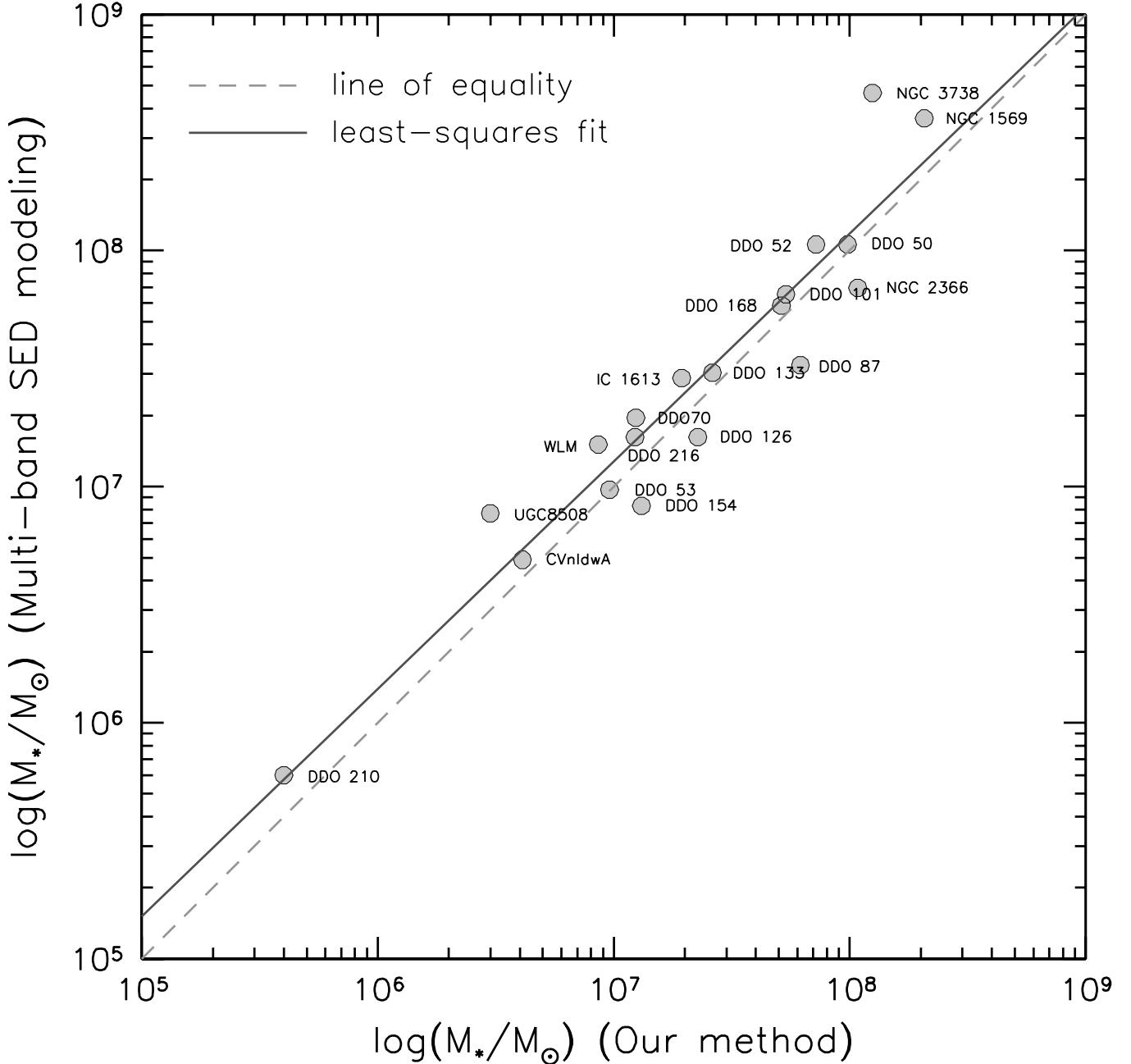


FIG. 1.— Comparison of the stellar masses derived using our method with those derived using the multi-band spectral energy distribution (SED) fitting technique described in Zhang et al. (2012). The solid and dashed lines indicate a least-squares fit with a slope of 0.96, and line of equality, respectively. The 1σ scatter of the fit to the data is 0.20 dex.

figures in the Appendix (e.g., Fig. A.3). In general, as discussed in Walter et al. (2008), the *Spitzer IRAC* $3.6\mu\text{m}$ image provides a pseudo dust free picture of old stellar populations in galaxies. However, unlike the case of the gas component whose mass can be directly estimated from HI observations, estimating the stellar mass in galaxies is critically dependent on the assumed stellar mass-to-light ratio (Υ_*) value which usually gives rise to the largest uncertainty when converting the luminosity profile to the mass density profile. In order to derive more reliable Υ_* values in the $3.6\mu\text{m}$ band, we use an empirical relation between galaxy optical colors and $\Upsilon_*^{3.6}$ values based on stellar population synthesis models (Bruzual & Charlot 2003; Bell & de Jong 2001) as given in Oh et al. (2008). Hereafter, we call these as model

$\Upsilon_*^{3.6}$ values. Using the derived $\Upsilon_*^{3.6}$ values, we convert the $3.6\mu\text{m}$ surface brightness profiles of our sample galaxies to stellar surface density profiles (see the figures in the Appendix, e.g., Fig. A.3).

As discussed in Oh et al. (2011a), the kinematic method combined with the model $\Upsilon_*^{3.6}$ values has been found to be reliable for estimating stellar masses of late-type dwarf galaxies based on a comparison of the derived stellar masses of the simulated dwarf galaxies to the input ones. This robustly supports the veracity of the methodology used for measuring the stellar masses of our sample galaxies. In addition, we also refer to the stellar masses of the sample galaxies derived using a spectral energy distribution (SED) fitting technique (Zhang et al. 2012). As shown in Fig. 1, the stellar masses derived using these

two independent methods show good agreement within the scatter.

Following Oh et al. (2011b), we calculate the corresponding rotation velocities of the stellar components of the sample galaxies from the derived surface density profiles. For this, we assume a vertical $\text{sech}^2(z)$ scale height distribution of stars with a ratio of $h/z_0=5$ where h and z_0 are the radial scale length and the vertical scale height of stellar disk in the $3.6\mu\text{m}$ surface brightness profiles, respectively (van der Kruit & Searle 1981; Kregel et al. 2002). Although it may overestimate the rotation velocities of the stellar components of galaxies with a fatter stellar disk, this is a valid assumption for most disk-dominated dwarf galaxies like our sample galaxies. The derived rotation velocities of the stellar components of the sample galaxies are shown in the figures in the Appendix (e.g., panel (d) of Fig. A.3), and the stellar masses estimated in this paper are given in Table 2.

The mean ratio of the masses between the gas and stellar components of the sample galaxies, $\langle M_{\text{gas}}/\text{M}_{\text{star}}^{\text{KIN}} \rangle$ is ~ 5.6 which is consistent with $\langle M_{\text{gas}}/\text{M}_{\text{star}}^{\text{SED}} \rangle \sim 7.2$ (where $\text{M}_{\text{star}}^{\text{SED}}$ is derived using a SED fitting technique) of 34 LITTLE THINGS sample galaxies in Zhang et al. (2012). This indicates that the majority of the baryons in our sample galaxies is in the form of gaseous components. As discussed earlier, given that the mass of gaseous components in galaxies can be reliably measured from HI observations without any critical assumption, the mass models of the baryons derived in this study are likely to provide a good description of the distribution of baryons in the galaxies. Therefore, any remaining uncertainties in the mass models of baryons are most likely to be insignificant and thus will not affect significantly the final mass models of the DM halos of our sample galaxies.

3.4. Disk-halo decomposition

In this Section, we decompose the total kinematics of the sample galaxies into the dynamical contributions of the baryonic disks and DM halos by disentangling the mass models of baryons from the total rotation curves. For this, we subtract in quadrature the rotation velocities of the gas and stellar components from the total rotation velocities, and obtain implied rotation curves for the DM halos as shown in the figures in the Appendix (e.g., see the left-lower panel of Fig. A.3). For a quantitative analysis of the DM distribution in the galaxies, we fit CDM (Navarro, Frenk & White 1996b, 1997; hereafter NFW) and spherical pseudo-isothermal halo models (e.g., Begeman et al. 1991), the two representative cusp- and core-like halo models, to the DM halo rotation curve, respectively. An Einasto profile which is a Sersic function in the context of CDM halos has been found to provide an equal or better fit to the halos in pure DM simulations compared to an NFW profile Navarro et al. (2004b). In this work, we use NFW profiles to quantify the DM halos of the sample galaxies for consistency with the previous DM mass modelling of THINGS and simulated dwarf galaxies in Oh et al. (2011a, 2011b). A relative comparison of the two halo models in terms of fit quality enables us to examine which model best describes the DM component of the sample galaxies, especially towards the centers of the galaxies.

3.4.1. Cusp-like halo model

The cosmologically motivated NFW halo model, the so-called ‘universal density profile’ which describes the cusp-like radial DM distribution found in DM-only Λ CDM simulations is given as,

$$\rho_{\text{NFW}}(R) = \frac{\rho_i}{(R/R_s)(1+R/R_s)^2}, \quad (6)$$

where ρ_i is correlated with the mean density of the Universe at the time of the collapse of the halo and R_s is the characteristic radius of the DM halo (Navarro et al. 1996b). This profile has been widely adopted to account for the DM distribution which steeply increases towards the centers of the halos in the simulations. This, a so-called cusp feature, can be well approximated by a power law, $\rho \sim R^\alpha$ with $\alpha \sim -1.0$ near the central region of the halos, giving a DM halo rotation velocity as follows,

$$V_{\text{NFW}}(R) = V_{200} \sqrt{\frac{\ln(1+cx) - cx/(1+cx)}{x[\ln(1+c) - c/(1+c)]}}, \quad (7)$$

where c is the concentration parameter defined as R_{200}/R_s . V_{200} is the rotation velocity at a radius R_{200} where the mass density contrast with the critical density of the Universe exceeds 200, and x is defined as R/R_{200} (Navarro et al. 1996b). In particular, the concentration parameter c is useful for quantifying the degree of DM concentration in galaxies. The measurement of the value of c in nearby galaxies provides an observational constraint on the central cusps predicted from Λ CDM simulations (McGaugh et al. 2007; see also de Blok et al. 2003).

3.4.2. Core-like halo model

As an alternative way to describe the DM distribution in a galaxy observationally motivated, spherical pseudo-isothermal halo models with a central constant-density core have been used in studies of galaxy rotation curves. The form of this core-like halo model is given as follows:

$$\rho_{\text{ISO}}(R) = \frac{\rho_0}{1 + (R/R_C)^2}, \quad (8)$$

where ρ_0 and R_C are the core-density and core-radius of a halo, respectively. This halo model is employed to describe the mass distribution of a DM halo with a sizeable constant density-core ($\rho \propto \rho_0$). Similarly, the corresponding rotation velocity to the pseudo-isothermal halo potential is given by,

$$V_{\text{ISO}}(R) = \sqrt{4\pi G \rho_0 R_C^2 \left[1 - \frac{R_C}{R} \tan\left(\frac{R}{R_C}\right) \right]}. \quad (9)$$

In order to quantify the DM distribution in the sample galaxies, we fit the model rotation velocities of the two halo models as given in Eqs. 7 and 9 to the kinematic DM signature of the galaxies. As shown in the figures in the Appendix (e.g., the left-lower panel of Fig. A.3), the core-like pseudo-isothermal halo models are mostly preferred over the cusp-like NFW models in terms of the quality of the fit (i.e., based on reduced χ^2 values). The fitted parameters of the halo models are presented in Table 2.

As given in Table 2, we derive the dynamical masses (M_{dyn}) of the sample galaxies using $V_{\text{ISO}}(R_{\text{max}})$ and

R_{\max} at which the outermost part of the rotation curve is measured. Given that rotation curves at the adopted R_{\max} are mostly rising in our sample galaxies, the derived dynamical mass, M_{dyn} with $V_{\text{ISO}}(R_{\max})$ and R_{\max} gives a lower limit of the halo mass. For reference, we also estimate the halo masses, M_{200} of our sample galaxies assuming an NFW halo model. However, the fitted values of c are unphysical (i.e., negative or close to zero) for a large fraction (11/26 galaxies) of the sample galaxies. We therefore derive a c value using an empirical relationship between $c - V_{200}$ from the WMAP²⁸ observations in McGaugh et al. (2007) by substituting V_{200} with the $V_{\text{ISO}}(R_{\max})$ adopted. We then fit the NFW halo model after fixing c and leaving only V_{200} as a free parameter, assuming a minimum disk (where the rotation curve is attributed to the DM halo only and the dynamical contribution of baryons is ignored). These c and V_{200} values are given in brackets in Table 2. Lastly, we derive the resulting halo mass, M_{200} with the newly estimated V_{200} using Eq. 3 in Oh et al. (2011a). As presented in Table 2, M_{200} values of the sample galaxies are larger than M_{dyn} values, which implies that our observations most likely do not reach the flat part of the rotation curves. This is consistent with the fact that the rotation curves of most sample galaxies are still rising at the last measured points.

Except in those few cases, such as DDO 70, DDO 101, DDO 154, DDO 210, DDO 216, and Haro 36 where CDM NFW halo models provide comparable fits to the DM rotation curves, the fitted values of the NFW halo parameters are unphysical (i.e., negative concentration parameter c , unphysically large values of V_{200}). This is consistent with results previously found in other nearby dwarf and LSB galaxies (e.g., de Blok & Bosma 2002; Kuzio de Naray et al. 2008; van Eymeren et al. 2009; Oh et al. 2011b). The slowly increasing DM rotation curves in the inner region of the sample galaxies reflect a halo whose gravitational potential is not deep enough to sustain the power-law DM density cusps that are as steep as $\rho \propto R^{-1.0}$. Hence, the cosmologically motivated cusp-like halo models are not able to adequately describe the observed solid-body rotation curves of our sample dwarf galaxies.

It may be argued that those galaxies that are equally well fitted by the two halo models indicate the possibility of a kinematic signature of central cusps in dwarf galaxies being consistent with ΛCDM simulations. Despite the high-resolution of LITTLE THINGS VLA H α observations, it is, however, most likely that the central regions of the galaxies are not fully sampled with a sufficient number of independent synthesized beams needed for distinguishing the inner steepness of the two halo models, accurately. As quantified in the parameter, $R_{\max}\text{H}\alpha_{\text{beam}}^{-1}$ in Table 2, this is mainly due to the small size of the rotating disk (e.g., DDO 210) or the relatively large distance (e.g., Haro 36). Given that the difference between the cusp- and core-like halo models is the most prominent in the central regions of galaxies, higher sampling of the inner regions of the galaxies is required before making a firmer conclusion on the signature of the potential central cusps. We will discuss this matter in a

more quantitative way in the following Section.

4. DARK MATTER DISTRIBUTION

In this Section, we compare the inner shape of rotation curves and DM density profiles of the sample galaxies with those of simulated dwarf galaxies from N-body+SPH ΛCDM simulations in order to examine their DM distributions near the centers. These comparisons between observations and simulations allow us to estimate the degree of cuspiness of the central DM distribution in a qualitative way. In addition, we also measure the inner DM density slopes of the galaxies to make a more quantitative comparison to the simulations.

4.1. Rotation curve shape

The characteristic shape of the steeply rising rotation curve inherent in the cusp-like DM distribution near the centers of simulated dwarf galaxies based on the ΛCDM paradigm can be used for a qualitative test of the simulations (Hayashi & Navarro 2006; see also Oh et al. 2011b, 2011a). This qualitative test is particularly useful in that a direct comparison between the observed and predicted DM rotation curves can be made without any additional assumption on the shape of the DM halo (e.g., a spherical or triaxial halo potential) which is needed for converting rotation curves to the corresponding density profiles, and the associated additional uncertainties this might introduce.

This is done in the left panel of Figs. 2, 3 and 4 where we scale the rotation curves of both our sample galaxies and ΛCDM NFW halos with respect to the rotation velocity $V_{0.3}$ at a radius $R_{0.3}$ which is where the logarithmic slope of the curve is $d\log V/d\log R = 0.3$ (see Hayashi & Navarro 2006). This enables us not only to make a relative comparison of the rotation curves between observations and simulations but also to accentuate the inner rotation curve shape which is sensitive to the degree of central DM concentration.

All the scaled (DM only) rotation curves of the 21 LITTLE THINGS sample galaxies for which *Spitzer* 3.6 μm image is available are overplotted in the upper-left panel of Fig. 5. We also overplot the median values of the rotation curves in each 0.1R/R_{0.3} bin. In the lower-left panel of Fig. 5, we also overplot the scaled rotation curves of seven dwarf galaxies from THINGS (three of them are also in LITTLE THINGS) as well as the two simulated dwarf galaxies presented in Governato et al. (2010) which all show a linear increase in their inner regions to the median values of the LITTLE THINGS rotation curves. In particular, the two simulated dwarf galaxies were affected by baryonic feedback processes (mainly repeated gas outflows driven by SN explosions) in such a way that the central cusps predicted from DM-only simulations are flattened (Governato et al. 2010). The flattened DM distribution results in slowly increasing rotation curves in the inner region of the simulated galaxies (see Oh et al. (2011a) for further discussion).

In line with the results in Oh et al. (2011a), the inner shape of the scaled rotation curves of the LITTLE THINGS sample galaxies falls mostly below that of the DM-only NFW halo models, indicating a shallower DM distribution near the centers. As shown in the lower-left panel of Figs. 2, 3 and 4, the discrepancy with NFW models becomes more pronounced when comparing these

²⁸ The Wilkinson Microwave Anisotropy Probe (Spergel et al. 2003; Spergel et al. 2007)

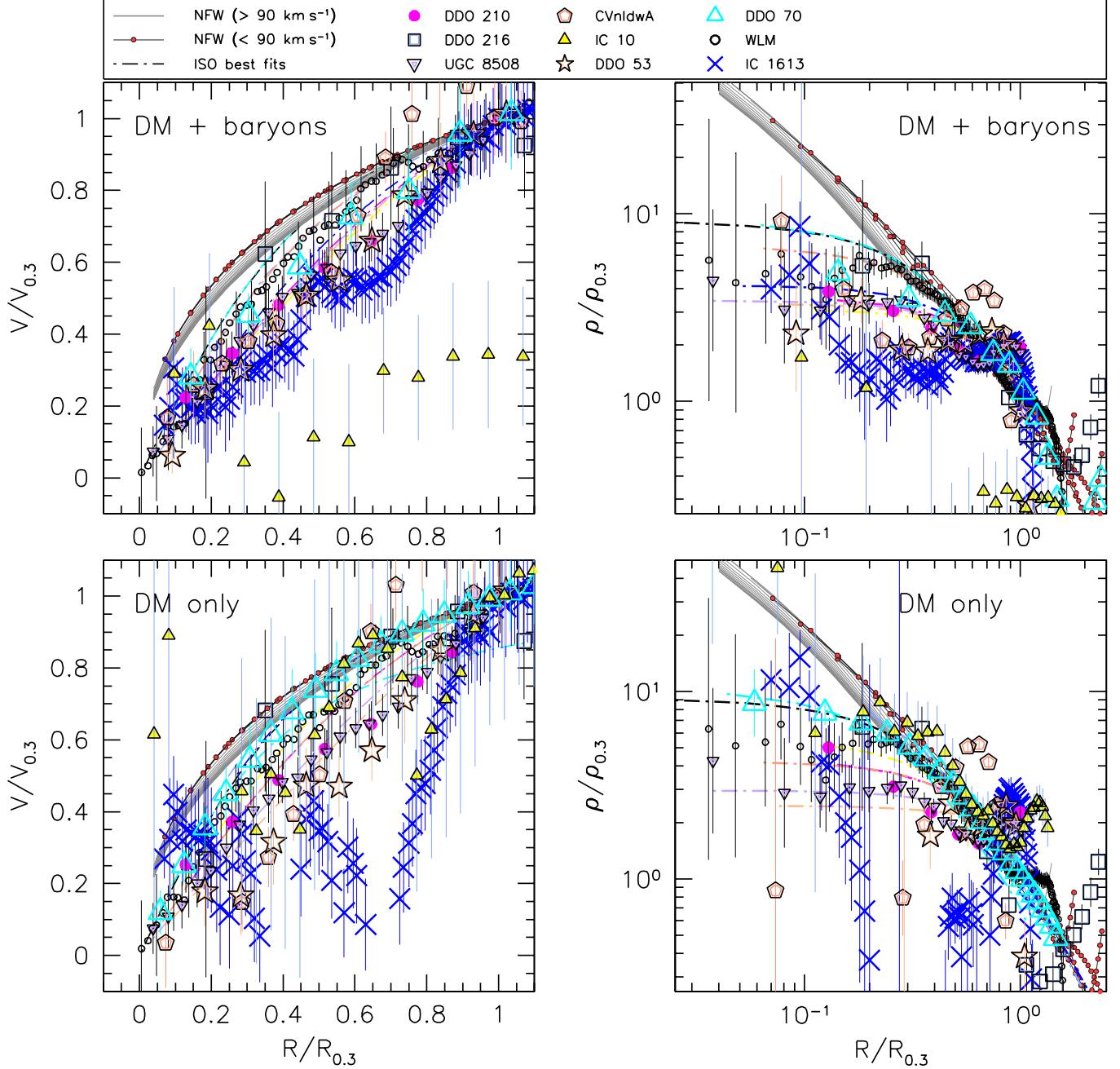


FIG. 2.— **Left panels:** The rotation curves of the first nine galaxies of the 26 LITTLE THINGS (in dynamical mass order) which are all scaled with respect to the rotation velocity $V_{0.3}$ at $R_{0.3}$ where the logarithmic slope of the rotation curve is $d\log V/d\log R = 0.3$ as described in Hayashi & Navarro (2006). The upper (DM + baryons) and lower (DM only) panels show the ones including and excluding the dynamical contribution by baryons, respectively. The grey solid and black solid lines with small dots indicate the CDM NFW dark matter rotation curves with V_{200} which is $> 90 \text{ km s}^{-1}$ and $< 90 \text{ km s}^{-1}$, respectively. The dashed lines (denoted as ISO) show the best fitted pseudo-isothermal halo models to the galaxies. **Right panels:** The corresponding dark matter density profiles derived using the scaled rotation curves in the left panels. The grey ($V_{200} > 90 \text{ km s}^{-1}$) and black solid lines with small dots ($V_{200} < 90 \text{ km s}^{-1}$) represent the CDM NFW models with the inner density slope $\alpha \sim -1.0$. The dashed lines indicate the best fitted pseudo-isothermal halo models with $\alpha \sim 0.0$. See Section 4 for more details.

models to DM rotation curves derived after subtracting the contribution from baryons from the total kinematics for each of the sample galaxies. Instead, they are more in line with those of both the THINGS dwarf galaxies and simulations (DG1 and DG2) where the effect of baryonic feedback processes is included. The linearly (or less steeply) rising rotation curves in the inner region of the galaxies indicate a nearly constant or shallower mass distribution towards the centers as found in the majority of nearby dwarf galaxies. The LITTLE THINGS sample

galaxies give no clear indication of the central cusps in their DM halos unlike the predictions from Λ CDM DM-only simulations.

4.2. Dark matter density profiles

As a more direct way to examine the central DM distribution of the sample galaxies and compare them with Λ CDM simulations, we derive their DM density profiles from the DM rotation curves decomposed in Section 3.4. A direct conversion of the rotation curve to the corre-

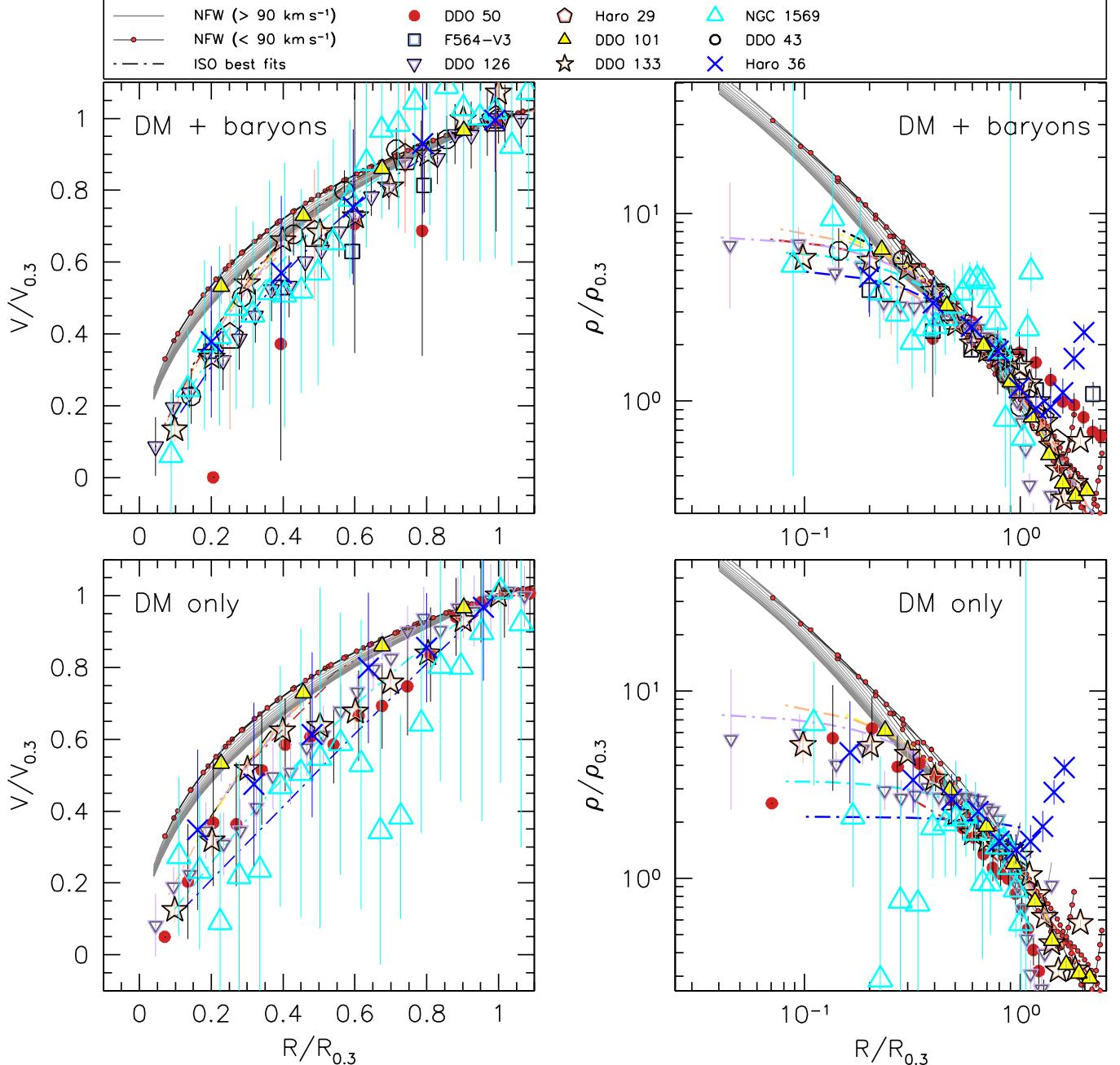


FIG. 3.— **Left panels:** The rotation curves of the other nine galaxies of the 26 LITTLE THINGS (in dynamical mass order) which are all scaled with respect to the rotation velocity $V_{0.3}$ at $R_{0.3}$ where the logarithmic slope of the rotation curve is $d\log V/d\log R = 0.3$ as described in Hayashi & Navarro (2006). The upper (DM + baryons) and lower (DM only) panels show the ones including and excluding the dynamical contribution by baryons, respectively. The grey solid and black solid lines with small dots indicate the CDM NFW dark matter rotation curves with V_{200} which is $> 90 \text{ km s}^{-1}$ and $< 90 \text{ km s}^{-1}$, respectively. The dashed lines (denoted as ISO) show the best fitted pseudo-isothermal halo models to the galaxies. **Right panels:** The corresponding dark matter density profiles derived using the scaled rotation curves in the left panels. The grey ($V_{200} > 90 \text{ km s}^{-1}$) and black solid lines with small dots ($V_{200} < 90 \text{ km s}^{-1}$) represent the CDM NFW models with the inner density slope $\alpha \sim -1.0$. The dashed lines indicate the best fitted pseudo-isothermal halo models with $\alpha \sim 0.0$. See Section 4 for more details (continued).

sponding DM density profile can be made by the following formula (see de Blok et al. 2001 for more details),

$$\rho(R) = \frac{1}{4\pi G} \left[2 \frac{V}{R} \frac{\partial V}{\partial R} + \left(\frac{V}{R} \right)^2 \right], \quad (10)$$

where V is a rotation velocity observed at a radius R , and G is the gravitational constant. For this conversion, we assume a spherical halo potential which is valid for most nearby galaxies (Trachternach et al. 2008). This

method has been used for deriving DM density profiles of dwarf and LSB disk galaxies, and proved to be reliable as found in the comparison of the derived DM density profiles of simulated dwarf galaxies with their input ones (Oh et al. 2011a).

In the right panel of Figs. 2, 3 and 4, we present the derived (dark) matter density profiles of the LITTLE THINGS sample galaxies. In addition, we calculate the median values of the DM density profiles as shown in the upper-right panel of Fig. 5. We also overplot the

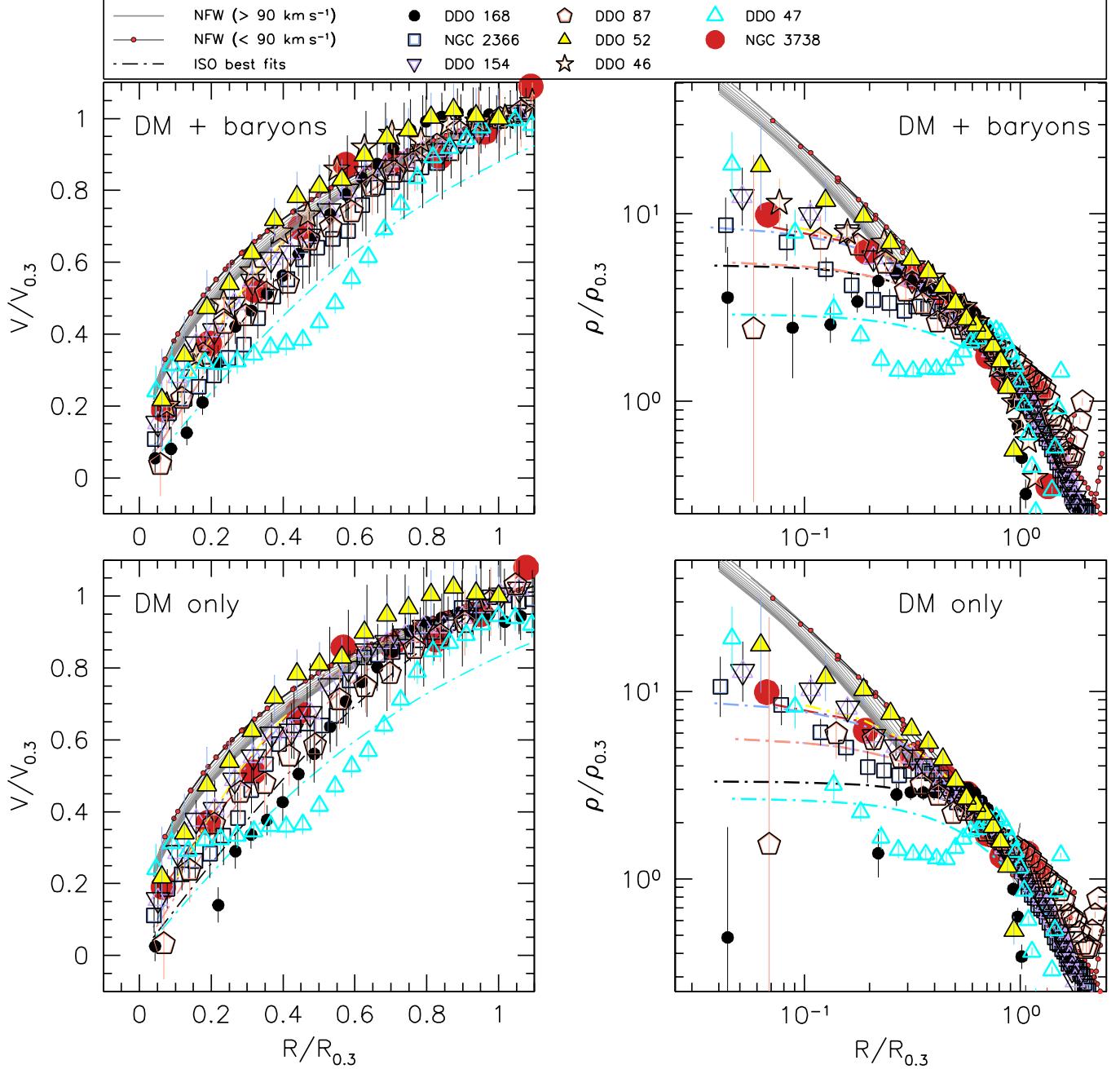


FIG. 4.— **Left panels:** The rotation curves of the remaining eight galaxies of the 26 LITTLE THINGS (in dynamical mass order) which are all scaled with respect to the rotation velocity $V_{0.3}$ at $R_{0.3}$ where the logarithmic slope of the rotation curve is $d\log V/d\log R = 0.3$ as described in Hayashi & Navarro (2006). The upper (DM + baryons) and lower (DM only) panels show the ones including and excluding the dynamical contribution by baryons, respectively. The grey solid and black solid lines with small dots indicate the CDM NFW dark matter rotation curves with V_{200} which is $> 90 \text{ km s}^{-1}$ and $< 90 \text{ km s}^{-1}$, respectively. The dashed lines (denoted as ISO) show the best fitted pseudo-isothermal halo models to the galaxies. **Right panels:** The corresponding dark matter density profiles derived using the scaled rotation curves in the left panels. The grey ($V_{200} > 90 \text{ km s}^{-1}$) and black solid lines with small dots ($V_{200} < 90 \text{ km s}^{-1}$) represent the CDM NFW models with the inner density slope $\alpha \sim -1.0$. The dashed lines indicate the best fitted pseudo-isothermal halo models with $\alpha \sim 0.0$. See Section 4 for more details (continued).

DM density profiles of the THINGS and simulated dwarf galaxies in the lower-right panel of Fig. 5. The radial fluctuation shown in some of the DM density profiles (e.g., IC 1613 etc.) is largely due to fluctuations in the derived rotation curves at the relevant radius. These are mainly because of either the effect of non-circular motions, the noise in velocity profiles with low S/N values, or both in the region.

As already implied by the solid-body like inner rotation curve shape of the LITTLE THINGS sample galaxies in

Section 4.1, their central DM density profiles are systematically shallower than those of the cusp-like DM density profiles predicted from DM-only Λ CDM simulations. As shown in the lower-right panel of Figs. 2, 3 and 4, the difference is even more pronounced in the comparison of the DM density profiles corrected for the baryons although the dynamical contribution by baryons is rather insignificant. The sample dIrr galaxies are more consistent with the core-like DM density profiles (dot-dashed lines in the right-hand frames of Figs. 2, 3 and 4) than NFW-type

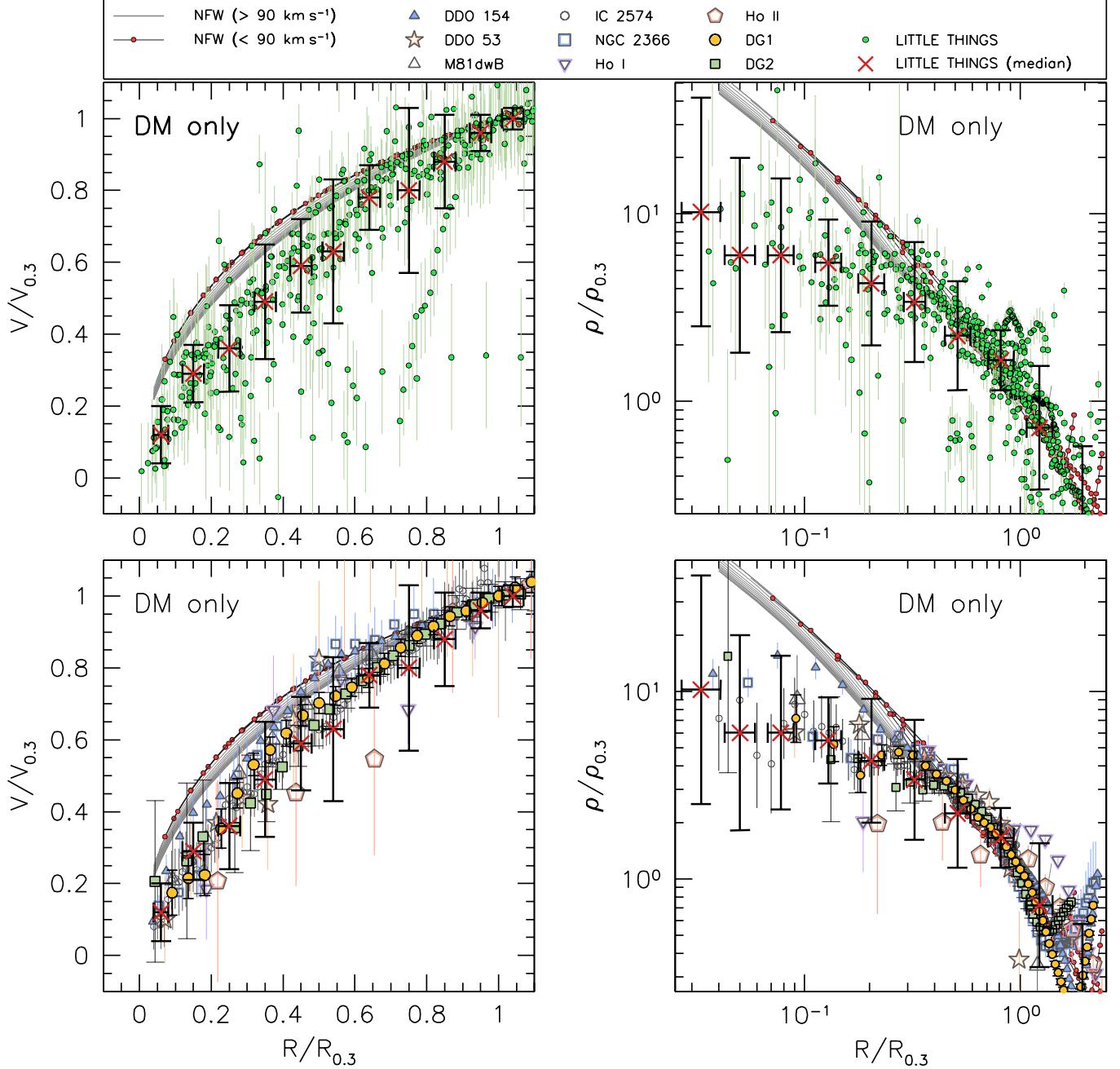


FIG. 5.— **Upper-left panel:** The (DM only) rotation curves (small dots) of the 21 LITTLE THINGS (including 3 THINGS galaxies) for which *Spitzer* 3.6 μ m image is available. These are all scaled with respect to the rotation velocity $V_{0.3}$ at $R_{0.3}$ where the logarithmic slope of the rotation curve is $d\log V/d\log R = 0.3$ as described in Hayashi & Navarro (2006). The ‘ \times ’ symbol represents the median values of the rotation curves in each $0.1R/R_{0.3}$ bin. The error bars show the 1σ scatter. **Lower-left panel:** The scaled rotation curves of the seven THINGS, and the two simulated dwarf galaxies (DG1 and DG2 in Governato et al. 2010) which are overplotted to the median values of the LITTLE THINGS rotation curves. The grey solid and black solid lines with small dots indicate the CDM NFW dark matter rotation curves with V_{200} which is > 90 km s⁻¹ and < 90 km s⁻¹, respectively. **Right panels:** The corresponding dark matter density profiles derived using the scaled rotation curves in the left panels. The grey ($V_{200} > 90$ km s⁻¹) and black solid lines with small dots ($V_{200} < 90$ km s⁻¹) represent the CDM NFW models with the inner density slope $\alpha \sim -1.0$. See Section 4 for more details.

profiles. This is much like the THINGS dwarf galaxies, and the simulated dwarfs (DG1 and DG2) with baryonic feedback processes as shown in the lower-right panel of Fig. 5.

We also measure the inner density slopes α of the DM density profiles to quantify the cuspiness of the central DM distribution. This yields a more quantitative comparison between the observations and simulations. As shown in the figures in the Appendix (e.g., panel (f) of

Fig. A.3), we perform a least squares fit (dotted lines) to the inner data points (grey dots) within a ‘break radius’. As described in de Blok & Bosma (2002; see also Oh et al. 2011b), we determine a break radius of a DM density profile where the slope changes most rapidly in the inner region of the profile. Following de Blok & Bosma (2002), we adopt the mean difference between the slopes which are measured including the first data point outside the break-radius and excluding the data point at the

break radius, respectively, as an errorbar $\Delta\alpha$ of the inner density slope. We measure the inner density slopes α , of the galaxies from their total matter (including both DM halo and baryons) as well as DM-only density profiles. The former, a so-called ‘minimum disk assumption’ that attributes the total rotation curve to the DM component only, gives a steeper inner density slope. Meanwhile, the latter where the dynamical contribution by baryons is subtracted from the total rotation curve allows us to examine the effect of the model $\Upsilon_*^{3.6}$ on the measured inner density slope. The mean values of the slopes of the 26 LITTLE THINGS dwarf galaxies are $\alpha_{\min} = -0.42 \pm 0.21$ and $\alpha_{\Upsilon_*,\text{pdf}} = -0.32 \pm 0.24^{29}$ for the minimum disk and the model $\Upsilon_*^{3.6}$ disk assumptions, respectively.

As expected, the slopes measured assuming the minimum disk are slightly steeper than those measured using the model $\Upsilon_*^{3.6}$ disk for most sample galaxies. However, the difference between the two slopes is largely insignificant since most of the sample galaxies are DM dominated as indicated by their low baryonic fraction. The measured logarithmic inner density slopes α of our sample galaxies are listed in Table 2.

The LITTLE THINGS sample dwarf galaxies do not in general agree with the steep logarithmic inner slope (~ 1.0) of the DM density profiles predicted from Λ CDM DM-only simulations. Instead, they show a range of shallower slopes being consistent with a core-like DM distribution at the centers, which supports the previous results found in nearby dwarf and LSB disk galaxies. Our results ($\alpha_{\min} = -0.42 \pm 0.21$; $\alpha_{\Upsilon_*,\text{pdf}} = -0.32 \pm 0.24$) are consistent with the mean logarithmic slope, $\alpha_{\min} = -0.29 \pm 0.07$, of the seven THINGS dwarf galaxies derived assuming a minimum disk in Oh et al. (2011b). Moreover, if we combine the sample dwarf galaxies from LITTLE THINGS and THINGS which have a similar data quality and whose inner density slopes are derived in exactly the same way, the mean value of the slopes of the 29 dwarf galaxies is $\alpha_{\min} = -0.40 \pm 0.24$. This shows good agreement within the error bars with $\alpha_{\min} = -0.2 \pm 0.2$ derived from LSB galaxies in de Blok et al. (2001) but a clear deviation from the $\alpha \sim -1.0$ predicted from Λ CDM DM-only simulations.

A few galaxies in our sample, such as DDO 101 and DDO 210 whose rotation curves are equally fitted by both NFW and pseudo-isothermal halo models appear to have relatively steeper slopes compared to the other ones. However, as discussed in Section 3.4, this could be due to insufficient sampling of the dark matter density profiles in the inner region. The gradient of the logarithmic density slope, $d\log\rho/d\log R$ gradually decreases towards the outer region of a galaxy, giving a steeper slope α . Therefore, the steeper a slope α is, the more data points in the outer regions are included when measuring the logarithmic slope of a DM density profile. This could, conceivably, be the case for galaxies with insufficient spatial sampling. As discussed above, the insufficient sampling mainly arises from either the smaller size of the HI disk (e.g., DDO 210) or the larger distances of the galaxies (e.g., Haro 36).

²⁹ The dynamical contribution by the stellar component is included in the DM density profiles of five galaxies where no *Spitzer* 3.6 μm image is available.

5. EFFECT OF SN FEEDBACK ON THE CENTRAL CU.PDF

As discussed in Governato et al. (2010), the constant-density cores observed near the centers of dwarf galaxies can be reconciled with simulations by taking the effect of baryonic feedback processes into account without a need for any explicit modification of the current Λ CDM paradigm. In particular, repeated gas outflows driven by SNe have been found to be efficient enough to redistribute the matter in dwarf galaxies, resulting in shallower DM density profiles as observed in nearby dwarf galaxies (Oh et al. 2011a).

The investigation of the effect of SN-driven gas outflows on the central DM distribution has been extended to low mass field dwarf galaxies using GASOLINE (Wadsley et al. 2004), a parallel SPH tree-code with multi-stepping (Governato et al. 2012; Pontzen & Governato 2012; Di Cintio et al. 2014; see also Pontzen & Governato 2014). In the simulations, the present-day stellar mass of galaxies ranges from $10^{9.8}$ down to $10^{4.5} M_\odot$ where the energy transfer from repeated gas outflows to the DM component becomes inefficient. More specifically, the SN feedback in small halos where less than 0.03% of the total amount of baryons is converted into stars is less effective at removing the central cusps and turning the cusp-like DM density profiles into core-like ones. According to this scenario, it is expected that the central DM distribution in these systems remains cuspy, and the inner slopes of their DM density profiles are steeper than those of higher mass counterparts (Governato et al. 2012; Zolotov et al. 2012). It implies that the central cusps predicted from Λ CDM DM-only simulations should survive in low mass halos. This underlines the cosmological importance of low mass dwarf galaxies in the local Universe for testing the ‘cusp/core’ problem.

As discussed earlier, the systematic uncertainties caused by low resolution radio observations are significantly reduced in the high-resolution HI data from LITTLE THINGS, which allows us to derive more accurate rotation curves of the sample galaxies and thus their central DM distributions. In addition, the *Spitzer* IRAC 3.6 μm data combined with model $\Upsilon_*^{3.6}$ values based on stellar population synthesis models and galaxy colors provide more reliable stellar masses of the sample galaxies.

In Fig. 6, we plot the inner density slopes α of the sample dwarf galaxies from both LITTLE THINGS and THINGS against their stellar masses M_* on a logarithmic scale. In addition, we also add those of a sample of THINGS disk galaxies whose mass models were derived in de Blok et al. (2008) in order to examine the $\alpha - M_*$ relationship in the higher mass regime. Compared to dwarf galaxies, the central kinematics of disk galaxies are usually dominated by a bulge component. It is therefore not trivial to perform a reliable disk-halo decomposition of the disk galaxies despite using the multi-wavelength data from THINGS whose data quality is comparable to that of LITTLE THINGS. For this reason, in Fig. 6, we use the inner density slopes of the disk galaxies measured assuming a minimum disk. As discussed in Section 4.2, it should be noted that the minimum disk assumption attributes the total kinematics of a galaxy to the DM only, ignoring the contribution of baryonic components, giving a lower limit on the inner density slope (i.e., a steeper

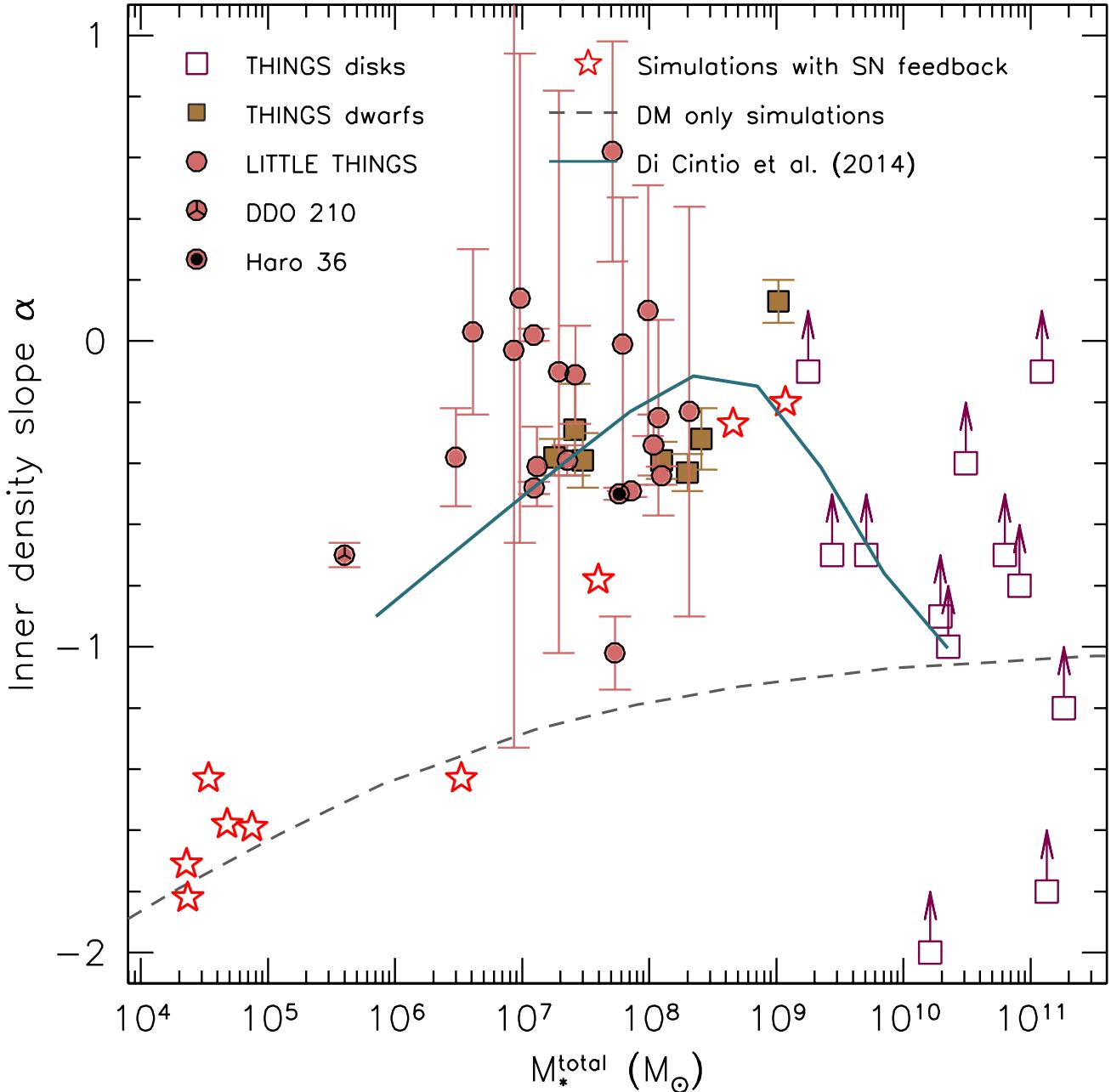


FIG. 6.— The inner dark matter density slope α of the sample galaxies from LITTLE THINGS (filled circles) and THINGS (filled squares, THINGS dwarfs; open squares, THINGS disk galaxies) against their total stellar masses. The dashed line indicates the α - M_*^{total} prediction from Λ CDM dark matter only simulations. The open stars and solid line represent the α vs. M_*^{total} of the resolved halos from the Λ CDM SPH simulations with baryonic feedback processes which are measured at 500 pc and $z=0$ (Governato et al. 2012; Di Cintio et al. 2014). See Section 5 for more discussions.

slope).

As shown in Fig. 6, the Λ CDM SPH simulations including the effect of baryonic feedback processes predict cusp-like DM distributions with steeper inner density slopes ($\alpha < -1.0$) in DM halos whose stellar mass is less than about $10^6 M_\odot$. As discussed in Governato et al. (2012), the lower the mass of the stellar component in a galaxy, the less the dynamical effect of SN feedback on the DM potential. Consequently, this results in that the central DM distribution in the low mass halo regime remains cuspy. According to this, the initial cusps formed in the early Universe would still exist today in low mass dwarf galaxies which have stellar masses less than $\sim 10^6 M_\odot$

where the repeated thermal energy injection from SN explosions becomes substantially inefficient, mainly due to the rapidly decreasing star formation efficiency in these systems (Governato et al. 2012; Pontzen & Governato 2012).

However, in Fig. 6, as the stellar mass of a galaxy increases, the inner DM density slope α becomes shallower in the Λ CDM SPH simulations including baryonic feedback processes (open stars) with respect to the counterpart in the DM-only simulations (dashed line). As discussed in Governato et al. (2010) and Pontzen & Governato (2012) (see also Di Cintio et al. 2014), the central DM cusps can be disrupted by the rapid gas injec-

tion caused by SN-driven gas outflows into the central region of galaxies, resulting in a shallower DM density distribution. This shows that DM-baryon interactions in dwarf galaxies through gas outflows play a critical role not only in forming bulgeless dwarf galaxies but also in flattening central cusps predicted from DM-only Λ CDM simulations. This demonstrates that proper modeling of DM-baryon interactions in hydrodynamical Λ CDM galaxy simulations is able to alleviate the long-standing tension associated with the central DM distribution in dwarf galaxies between simulations and observations. We refer to Governato et al. (2012) (see also Pontzen & Governato 2012; 2014 and Di Cintio et al. 2014) for a detailed discussion of the effect of SN feedback on the central DM distribution in dwarf galaxies.

Meanwhile, in Fig. 6, the trend of slope change $d\log\alpha/d\log M_*$ predicted from dwarf galaxy simulations with baryonic feedback processes is reversed in massive disk galaxies (open boxes) where a bulge component becomes dominant in the central kinematics. As discussed in Di Cintio et al. (2014), the effect of SN feedback can be surpassed by the deep gravitational potential which is caused by the bulge component in the central region. However, as discussed earlier, the steep slopes of the THINGS disk galaxies could be partially affected by the minimum disk assumption used for deriving their DM density profiles. The central kinematics of a disk galaxy with a substantial bulge component is sensitive to even small uncertainties in Υ_* when converting the luminosity profile to the mass density profile. This makes it difficult to perform an accurate disk-halo decomposition of bulge-dominated disk galaxies. This stresses the usefulness of bulge-less dwarf galaxies in testing the effect of SN feedback on the central DM distribution of galaxies.

As shown in Fig. 6, the simulated dwarf galaxies which have comparable stellar masses ranging from 10^6 to $10^9 M_\odot$ show good agreement with the majority of the sample dwarf galaxies from THINGS and LITTLE THINGS. However, as already discussed in Section 4.2, the rotation curves of some LITTLE THINGS sample galaxies (e.g., DDO 70, DDO 101, DDO 154, DDO 210 and Haro 36) are equally well fitted by CDM NFW and pseudo-isothermal halo models in terms of χ^2_{red} values. Moreover some of them, such as DDO 101, DDO 210 and Haro 36, appear to have relatively steep inner density slopes (although Haro 36 is defined as a blue compact dwarf galaxy which usually shows a steep increase in rotation velocity in the inner region mainly due to young starburst components formed during galaxy interaction or mergers; e.g., Bekki 2008). In particular, DDO 210 is a good candidate for testing the SN feedback efficiency scenario by Governato et al. (2012) given that its stellar mass falls within the regime where according to the simulations primordial CDM cusps are expected to survive. In Fig. 6, DDO 210 shows no distinct signature of the central cusp ($\alpha \sim -1.0$) given its corresponding stellar mass, deviating from the prediction of Λ CDM DM-only simulations (dashed line). Nevertheless, in a qualitative sense, its relatively steep inner density slope $\alpha \sim -0.70$ compared to the others still leaves a room for a potential signature of the central cusp, which supports the lower SN energy injection scenario in low mass dwarf galaxies. However, as noted earlier, the resolution of LITTLE THINGS HI observations is not high enough to resolve the small HI

disk ($\sim 60''$ diameter) and distinguish between cusp- and core-like DM behaviour near the center of DDO 210. We note that DDO 210 has the most compact HI disk (in terms of the beam size, e.g., $R_{\max} \text{HI}_{\text{beam}}^{-1} \sim 6.0$ as given in Table 2) of the sample galaxies. Likewise, some of our sample galaxies may still suffer from beam smearing which is discussed in the following section.

6. INNER DENSITY SLOPE VS. RESOLUTION

As in de Blok et al. (2001), for a quantitative examination of the beam smearing effect on our sample galaxies, we plot the inner density slopes α of the galaxies including the THINGS sample as well as the two simulated dwarf galaxies (DG1 and DG2) modelled by Governato et al. (2010) against the observed radii of their innermost point R_{inner} in Fig. 7. For the sample dwarf galaxies from LITTLE THINGS, we use the slopes derived assuming the model $\Upsilon_*^{3.6}$ disk. We also show the α - R_{inner} relations of the NFW and pseudo-isothermal halo models as solid and dotted lines, respectively, derived using their analytical formulas as given in Eqs. 6 and 8.

As shown in Fig. 7, most sample galaxies show significant deviations from the predicted α - R_{inner} trend (solid line) of Λ CDM NFW halos at around a R_{inner} of ~ 0.2 kpc. Instead, they are more consistent with those of pseudo-isothermal halo models with different core-radii (dotted lines) as well as the earlier results found from LSB galaxies (grey symbols) in de Blok & Bosma (2002). However, the clear difference between the two halo models (i.e., NFW and pseudo-isothermal) at high resolutions (e.g., $R_{\text{inner}} < 0.5$ kpc) becomes ambiguous as the innermost radius R_{inner} of a given DM density profile increases. For example, a galaxy with a larger R_{inner} (i.e., low resolution) tends to show a steeper inner slope of DM density profile. The larger R_{inner} makes it lie in the region where the slopes of the two halo models are approximately similar to each other. In addition, as discussed in Section 4.2, the derivative $d\log\rho/d\log R$ of a DM density profile on a logarithmic scale decreases towards the outer region of a galaxy. If the DM density profile is affected by beam smearing, the break radius of the profile which is determined when measuring the inner slope tends to migrate into the outer regime where $d\log\rho/d\log R$ has a lower value. Therefore, the inner density slope α within the break radius is most likely to be steeper than the ones derived from well sampled profiles. This could be the case of DDO 101 and DDO 210. In particular, DDO 101 is most likely to be affected by the beam smearing effect as shown in Fig. 7.

Yet higher resolution velocity fields obtained with radio interferometers or using other tracers such as integral field mapping are required to study the effect of SN feedback on the central cusps of the lowest mass dwarf galaxies. Such high-resolution observations of low mass dwarf galaxies would provide an ultimate test of the Λ CDM paradigm. Unlike clusters of galaxies where the depth of the gravitational potential well is deep enough to retain warm DM (WDM) as well as CDM, there is no room for WDM in dwarf galaxies inhabiting DM halos with much shallower potential wells. Therefore, finding a signature of a central cusp in dwarf galaxies will prove that there is at least some cold DM in the Universe. This again highlights the cosmological importance of low mass dwarf galaxies, not only for resolving the ‘cusp/core’ contro-

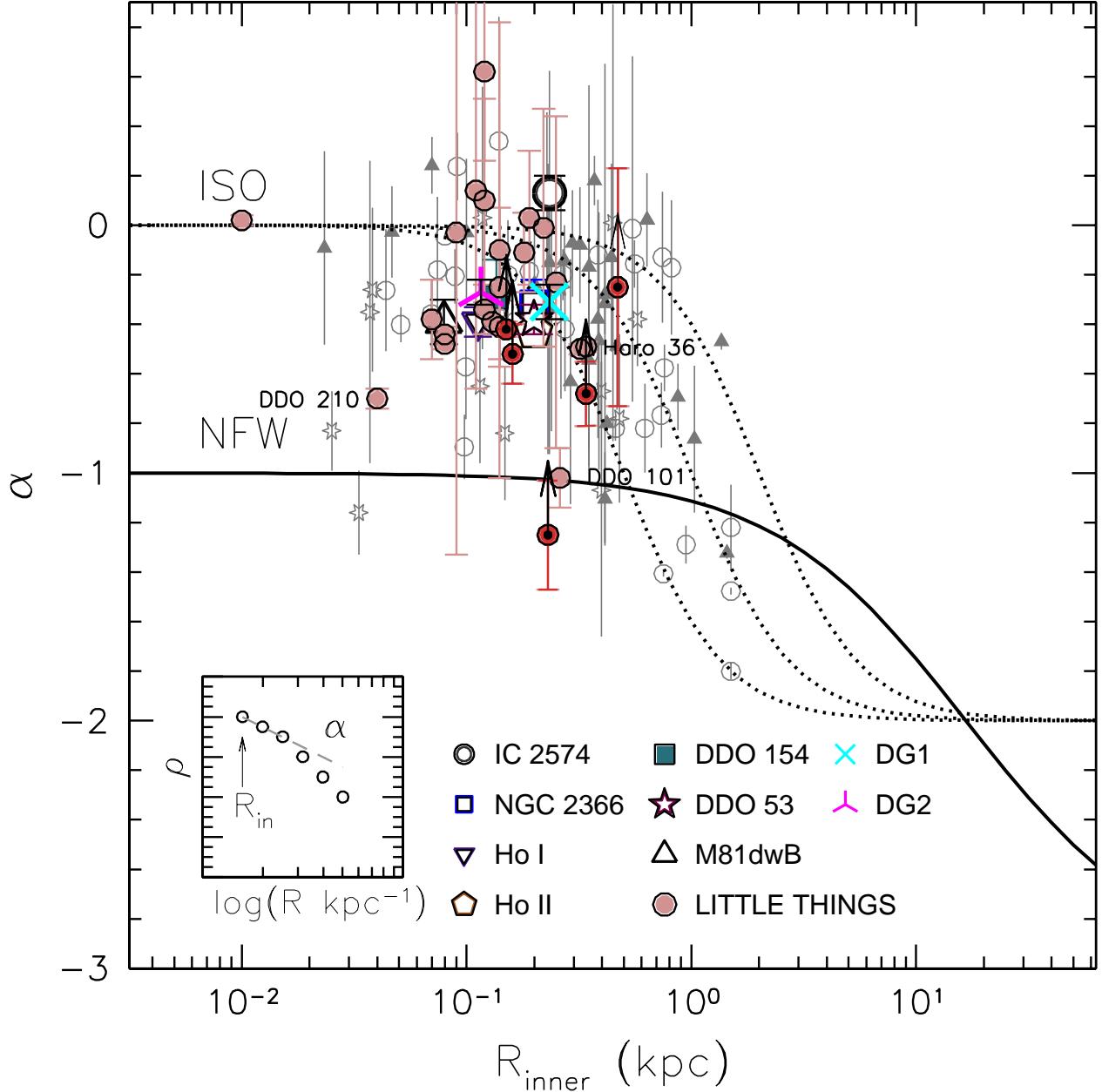


FIG. 7.— The inner slope of the dark matter density profiles α vs. the radius R_{in} of the innermost point within which α is measured as described in the small figure (de Blok et al. 2001). The α - R_{in} of the sample galaxies from LITTLE THINGS, THINGS and the two simulated dwarf galaxies (DG1 and DG2: Governato et al. 2010) as well as the previous measurements (grey symbols) of LSB galaxies (open circles: de Blok et al. 2001; triangles: de Blok & Bosma 2002; open stars: Swaters et al. 2003). Filled circles with arrows indicate the galaxies of which inner density slopes are measured assuming a ‘minimum disk’, giving a steeper slope. The solid and dotted lines represent the α - R_{in} trends of dark-matter-only Λ CDM NFW and pseudo-isothermal halo models, respectively. See Section 6 for more details.

versy in Λ CDM simulations but also as an indirect proof for the existence of CDM in the Universe.

7. CONCLUSION

In this paper we derive the rotation curves of 26 dwarf galaxies culled from LITTLE THINGS, and examine their DM distributions near the centers of the galaxies. From this, we address the ‘cusp/core’ problem which has been one of the long-standing problems in Λ CDM simulations on galactic scales. The high-resolution LITTLE THINGS HI data ($\sim 6''$ angular; $\sim 2.6 \text{ km s}^{-1}$ spectral) complemented with optical and *Spitzer* IRAC $3.6\mu\text{m}$ images are sufficiently detailed to resolve the central region of the sample galaxies where the cusp- and core-like halo

models are clearly distinguished.

In particular, we use the bulk velocity fields of the galaxies extracted using the method described in Oh et al. (2008) to correct for turbulent random non-circular gas motions. This enables us to derive more reliable rotation curves and thus more accurate DM distributions in the galaxies. We corrected for the modest dynamical contribution by baryons in dwarf galaxies by using *Spitzer* IRAC $3.6\mu\text{m}$ images combined with model $\Upsilon_*^{3.6}$ values based on stellar population synthesis models. This allowed us to derive robust mass models of the stellar components of the galaxies and thus better constrain their central DM distributions.

From this, we found that the decomposed DM rotation

curves of most sample galaxies are well matched in shape to those of core-like halos which are characterised by a linear increase of rotation velocity in the inner region. We also derive the DM density profiles of the sample galaxies and quantify the degree of the central DM concentration by measuring the logarithmic inner slopes of the profiles. The mean value of the inner slopes α of the 26 sample galaxies is -0.32 , which indicates a mass distribution with a sizable constant density-core towards the centers of the galaxies. This is consistent with that found in most nearby dwarf galaxies (e.g., LSB galaxies in de Blok & Bosma 2002; THINGS dwarfs in Oh et al. 2011b) which all show a linear increase in the inner shapes of their rotation curves resulting in shallower inner density slopes ($<\alpha> \sim -0.2$). Considering the fact that observational uncertainties are significantly reduced in the high-resolution LITTLE THINGS data, the core-like DM distribution found in our sample galaxies provides a stringent observational constraint on the central DM distribution of halos in Λ CDM simulations.

We find that the derived slopes of the DM density profiles do not agree with the cusp predicted by Λ CDM DM-only simulations. However, recent cosmological N-body SPH galaxy simulations by Governato et al. (2010) (see also Governato et al. 2012; Pontzen & Governato 2012) have shown that the discrepancy between observations and simulations can be reconciled within the Λ CDM paradigm by considering the dynamical effect of baryonic feedback processes on the central cusps. According to the simulations, DM-baryon interactions in dwarf galaxies through gas outflows driven by SN explosions play a critical role not only in forming bulgeless dwarf galaxies but also in turning central cusps into cores.

As discussed in Oh et al. (2011a), the slowly rising rotation curves and the resulting shallower DM density profiles of the simulated dwarf galaxies with SN feedback are qualitatively similar to those of the dwarf galaxies from THINGS. This is also the case for our sample galaxies from LITTLE THINGS whose DM rotation curves in the inner region rise too slowly to match the steep rotation curves of CDM halos. This shows that proper modelling of DM-baryon interactions in Λ CDM galaxy simulations is able to alleviate the long-standing tension between observations and simulations regarding the central DM distribution in dwarf galaxies.

Notwithstanding the dominant trend of core-like DM distribution in the LITTLE THINGS sample galaxies, some of the sample galaxies, such as DDO 210 and Haro 36 are equally well fitted by core- and cusp-like halo models in describing their DM rotation curves. They have relatively steeper inner density slopes with $\alpha \approx -0.70$ and -0.50 for DDO 210 and Haro 36, respectively, compared to the mean value (DM only) of the rest of the sample (~ -0.29). It is possible, however, despite the high angular resolution of the LITTLE THINGS H α data, the inner density slopes of these two galaxies are affected by beam smearing (the H α disk of DDO 210 being intrinsically small and Haro 36 being one of the more distant dwarfs). Insufficient spatial resolution in the inner region

of the galaxies acts to results in steeper observed inner DM density slopes.

According to the latest N-body SPH simulations of dwarf galaxies with baryonic feedback processes (Governato et al. 2012), the SN feedback in low mass dwarf galaxies with a stellar mass less than $10^6 M_\odot$ is not sufficient to disrupt the central cusps, the repeated energy injection from SN explosions into DM halos becoming inefficient, largely due to low star formation efficiencies in these low mass systems. However, previous observational studies regarding the ‘cusp/core’ problem have mostly focused on relatively massive dwarf galaxies for which reliable rotation curves are available. Low mass dwarf galaxies have been usually excluded for the study of the central DM distribution due to the low amplitude of their maximum rotation velocities. The kinematics of such low mass dwarf galaxies is more vulnerable to kinematic disturbances like non-circular motions in galaxies compared to more massive ones, and also more sensitive to additional corrections (e.g., asymmetric drift) made when deriving rotation curves. In this respect, the possibility of a selection effect in favor of relatively massive dwarf galaxies where the effect of SN feedback is enough to turn the central cusps into cores should be considered.

It would therefore be worthwhile to perform high-resolution follow-up observations, for example using an optical integral field unit, of low mass dwarf galaxies including some of the LITTLE THINGS dwarf galaxies (e.g., DDO 210) whose stellar masses lie in the regime where primordial CDM cusps are predicted from the simulations (Governato et al. 2012). These high-resolution observations will enable us to achieve a finer sampling of the central region of the galaxies, and thus more accurate inner DM density profiles. From this, more stringent observational constraints on the central cusp of low mass dwarf galaxies could be provided. Moreover, an accurate measurement of the DM distribution in these low mass dwarf galaxies will provide an ultimate test for the CDM paradigm given that dwarf galaxies inhabiting DM halos with a shallow potential well have only room for CDM, unlike clusters of galaxies whose gravitational potential is deep enough to retain WDM as well as CDM. Therefore, the presence or absence of a signature of the central cusp in these low mass halos will provide a critical observational test, either supporting or falsifying the Λ CDM paradigm.

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Appendix. DATA AND KINEMATIC ANALYSIS

In this Appendix, we present the data and kinematic analysis of the 26 LITTLE THINGS (Hunter et al. 2012) dwarf galaxies. For each galaxy, we show the (1) data, (2) kinematic analysis, and (3) mass modelling with descriptions:

A. Data— **(a)** Integrated HI intensity map (moment 0). The contour levels start at $+3\sigma$ in steps of $+3\sigma$. **(b)** Intensity-weighted mean velocity field (moment 1). **(c)** Velocity dispersion map (moment 2). **(d)** *Spitzer* IRAC 3.6 μ m image obtained from the archives including ‘*Spitzer* Infrared Nearby Galaxies Survey’ (SINGS; Kennicutt et al. 2003) and *Spitzer* ‘Local Volume Legacy’ (LVL; Dale et al. 2009). **(e)** Bulk velocity field extracted using the method described in Oh et al. (2008). **(f)** Velocity field of strong non-circular motions as in Oh et al. (2008). **(g)** Model velocity field of the tilted-ring model derived using the bulk velocity field in the panel (e). **(h)** Velocity field of weak non-circular motions as in Oh et al. (2008). **(i), (j)** Position-velocity diagram taken along the average position angle of the major and minor axes as given in Table 1. The dashed lines indicate the systemic velocity and position of the kinematic center derived in this paper. The bulk (black dots) and asymmetric drift corrected bulk (yellow) rotation curves are overplotted. The curves are converted back to radial velocities using the geometrical parameters determined by the tilted-ring analysis as listed in Table 1. Moment maps and velocity fields are extracted using the robust-weighted data cubes. The beam size is indicated by the ellipse in the bottom-right corner of each panel. See Hunter et al. (2012) for a detailed description of the data cube.

B. Kinematic analysis

• **Rotation curves**— The tilted-ring model derived using the bulk velocity field as given in the panel **(e)** of **A. Data**. The open grey circles shown in all panels are the fit results with all ring parameters (i.e., XPOS, YPOS, VSYS, PA, INCL and VROT) free. The grey filled dots in the VROT panel indicate the final rotation velocity derived using the entire velocity field after fixing other ring parameters to the values as shown in other panels (solid lines). The upright and upside-down triangles show the rotation velocities derived using the receding and approaching sides, respectively, while keeping other ring parameters at the values indicated in the other panels with solid lines.

• **Asymmetric drift correction**— **(a)** grey dots represent the radial asymmetric drift correction, σ_D . The open circles indicate the derived rotation velocity from the tilted-ring analysis, and the black dots show the one corrected for asymmetric drift. **(b)** Azimuthally averaged HI velocity dispersion. **(c)** Radial HI surface density derived applying the kinematic geometry from the tilted-ring analysis. **(d)** The dashed line shows the fit of the analytical function given in Eq. 4 to $\Sigma_{\text{HI}}\sigma^2$. See Section 3.1.2 for more details.

• **Harmonic analysis**— Harmonic decompositions of the bulk (black dots) and intensity-weighted mean (moment 1) velocity fields. The decomposition is made using the RESWRI task in GIPSY. The solid and dashed lines in the middle-right panel indicate average global elongations of the halo potential as described in Schoenmakers et al. (1997; 1999) derived using the bulk and moment 1 velocity fields, respectively. The median absolute amplitudes ($\langle A \rangle$) and the phases (ϕ_1 , ϕ_2 , and ϕ_3) of each component are presented in the bottom panels.

C. Mass modelling

• **Mass models of baryons**— **(a)** *Spitzer* IRAC 3.6 μ m surface brightness profile derived applying the kinematic geometry from the tilted-ring analysis. **(b)** The stellar mass-to-light value in the 3.6 μ m, $\Upsilon_*^{3.6}$ derived using the empirical relation described in Oh et al. (2008) which is based on stellar population synthesis models in Bruzual & Charlot (2003) and Bell & de Jong (2001). We refer to Oh et al. (2008) for a full description. **(c)** Stellar mass surface density profile in the 3.6 μ m. **(d)** The rotation velocity for the stellar component derived using the mass density profile in panel (c). **(e)** Azimuthally averaged optical color, $B - V$. **(f)** Gas column density derived applying the kinematic geometry from the tilted-ring analysis. **(g)** Mass surface density profile for the gas component which is scaled up by 1.4 to account for Helium and metals. **(h)** The resulting rotation velocity for the gas component.

• **Disk-halo decomposition**— The grey dots indicate the bulk rotation curve which is corrected for asymmetric drift. The dotted and dash-dotted lines represent the rotation velocities of the gas and stellar components, respectively. The open circles show the rotation velocity of dark matter halo only, derived after subtracting the contribution to the rotational velocity of the baryons. The dashed and solid lines are the fits of Λ CDM NFW and pseudo-isothermal halo models to the dark matter-only curve (open circles). The reduced χ^2 value for each halo model is denoted. The dots (DM rotation curve – NFW model) and open circles (DM rotation curve – pseudo-isothermal model) in the lower panel represent the velocity differences between the rotation curve of the dark matter halo and the best fit halo models.

• **Mass density profile**— The circles indicate the mass density profile derived from the asymmetric drift corrected

bulk rotation curve which includes the dynamical contributions by dark matter halo and baryons (i.e., minimum disk assumption). The squares show the one derived from the dark matter rotation curve where the dynamical contribution by baryons is subtracted (dark matter only). The logarithmic inner slope α of the dark matter density profile is measured by a least squares fit (dotted line) to the inner data points (filled squares). The dashed and solid lines show the density profiles derived from the best fit NFW and pseudo-isothermal halo models, respectively.

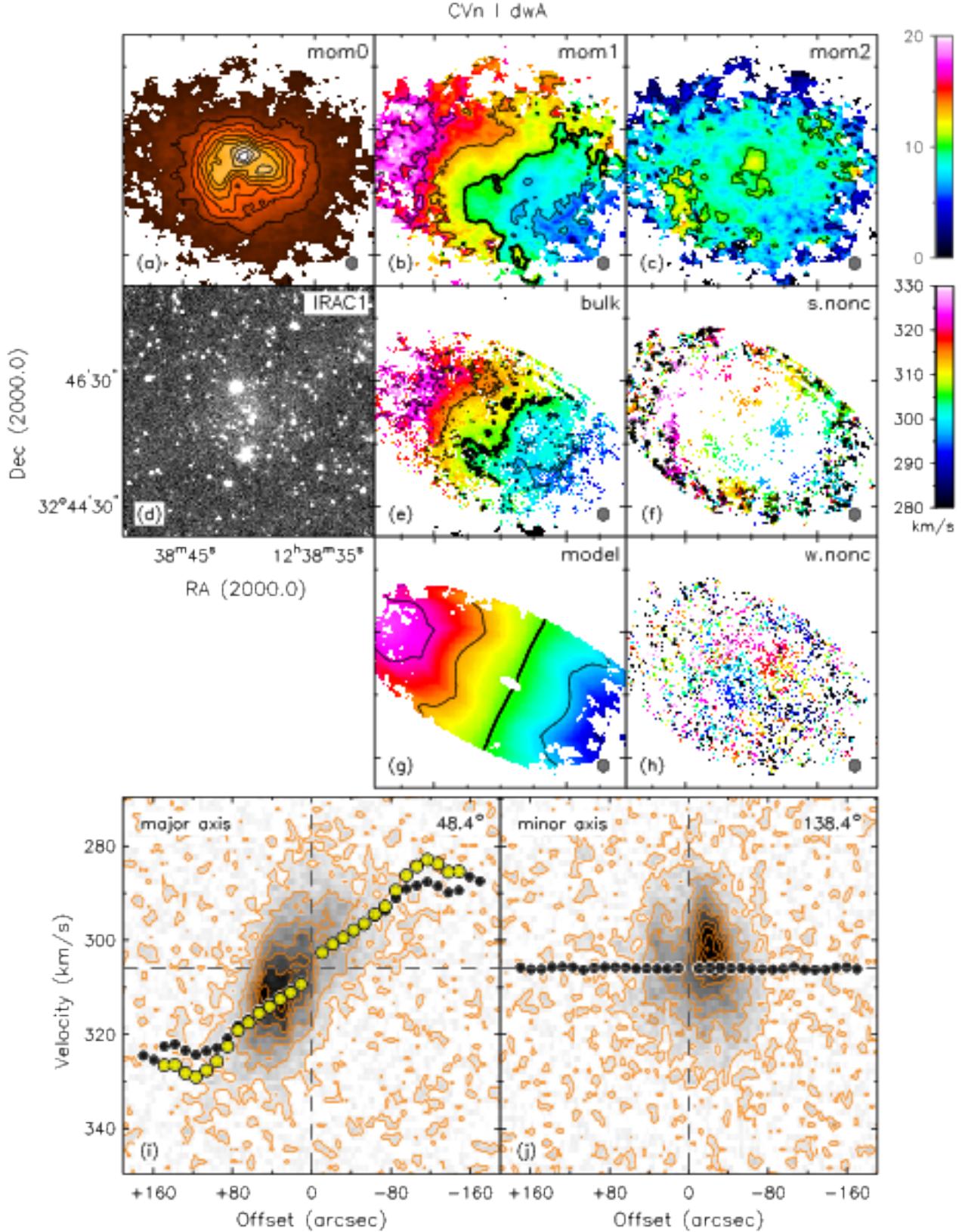


FIG. A.1.— HI data and *Spitzer* IRAC 3.6 μm image of CVnIdwA. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s⁻¹. Velocity dispersion contours run from 0 to 20 km s⁻¹ with a spacing of 5 km s⁻¹. See Appendix section A for details.

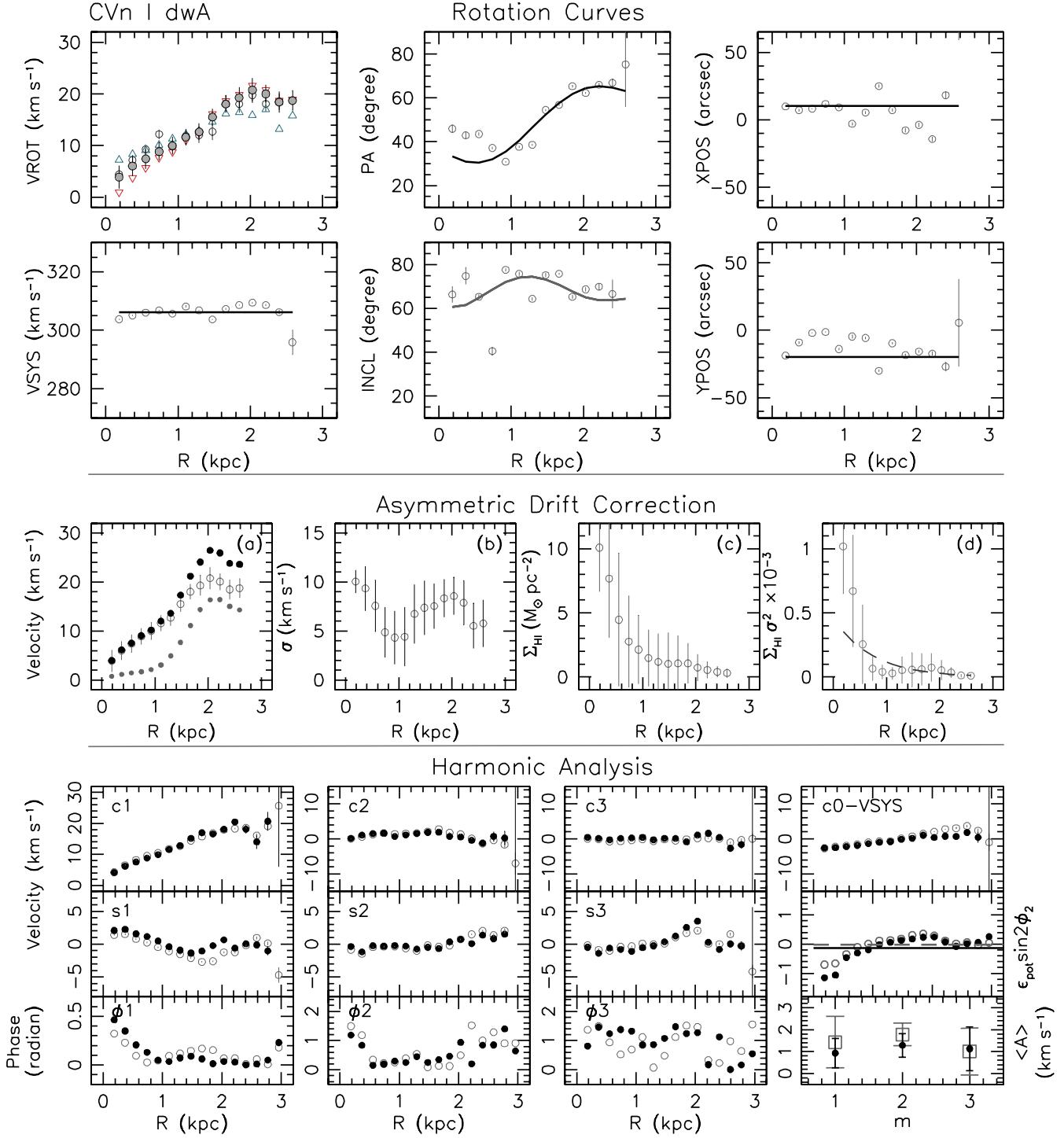


FIG. A.2.— Rotation curves, asymmetric drift correction and harmonic analysis of CVnIdwA. See Appendix section B for details.

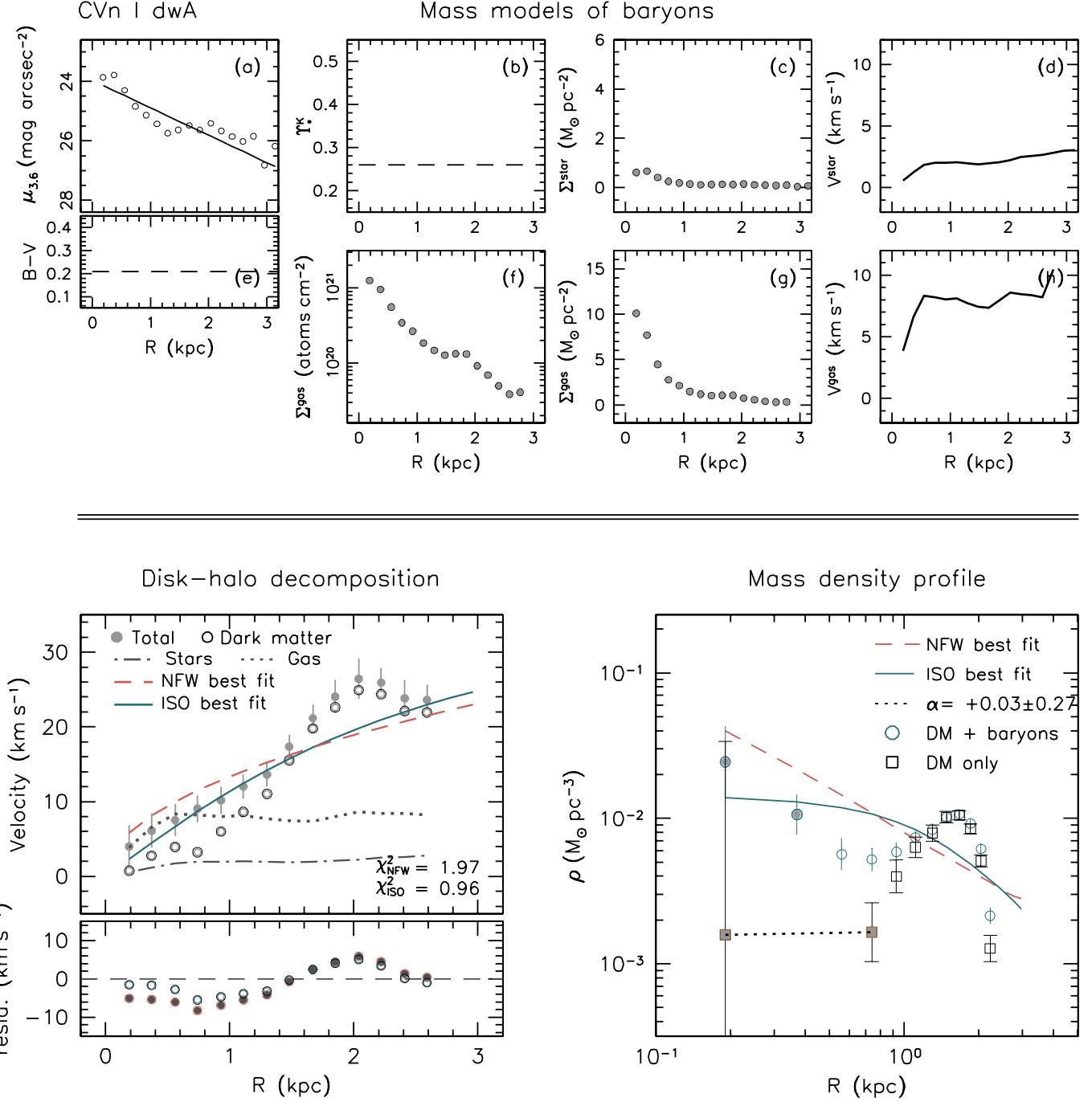


FIG. A.3.— The mass models of baryons, disk-halo decomposition and mass density profile of CVnIdwA. Please refer to the text in Sections 3 and 4 for full information.

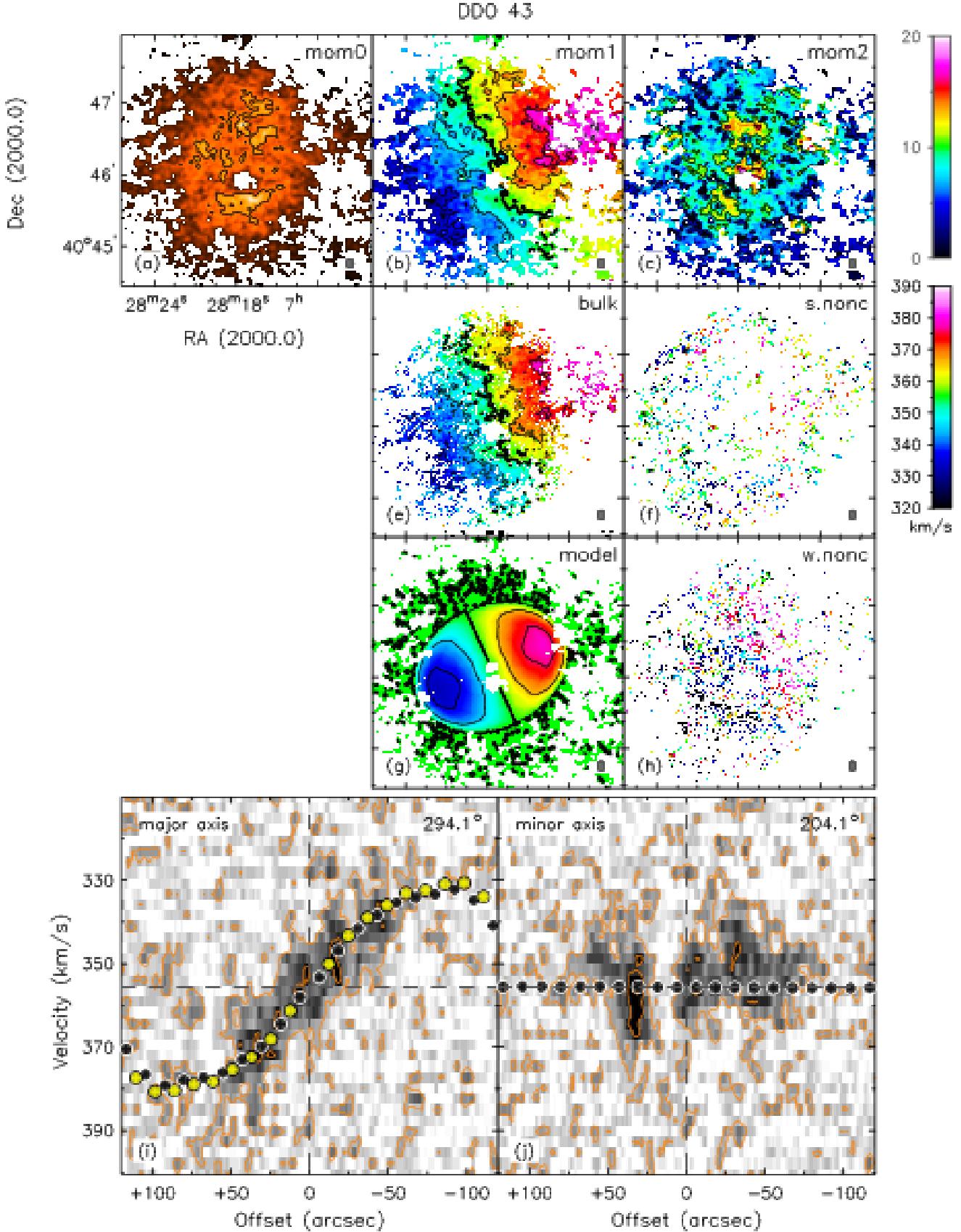


FIG. A.4.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 43. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

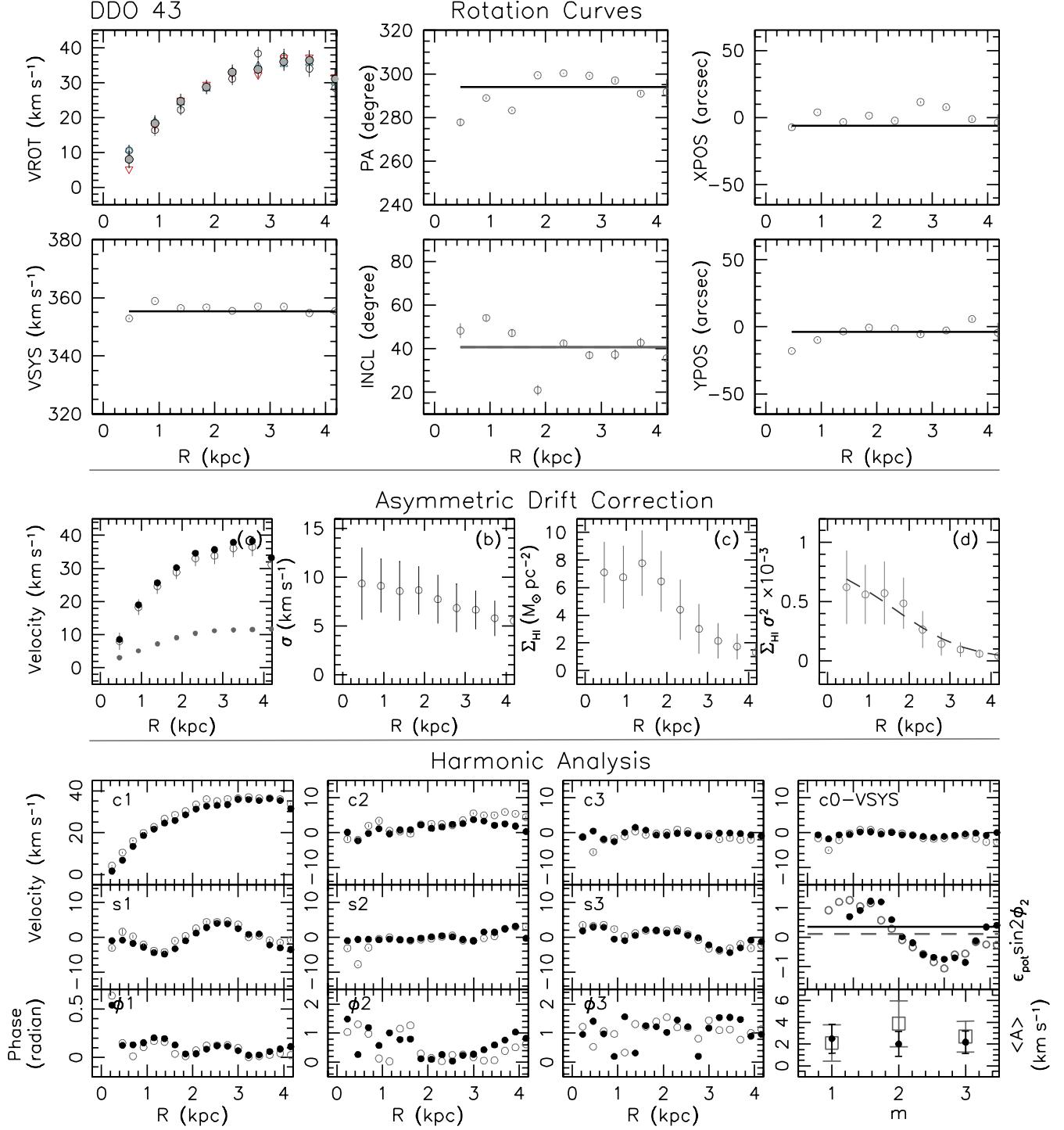


FIG. A.5.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 43. See Appendix section B for details.

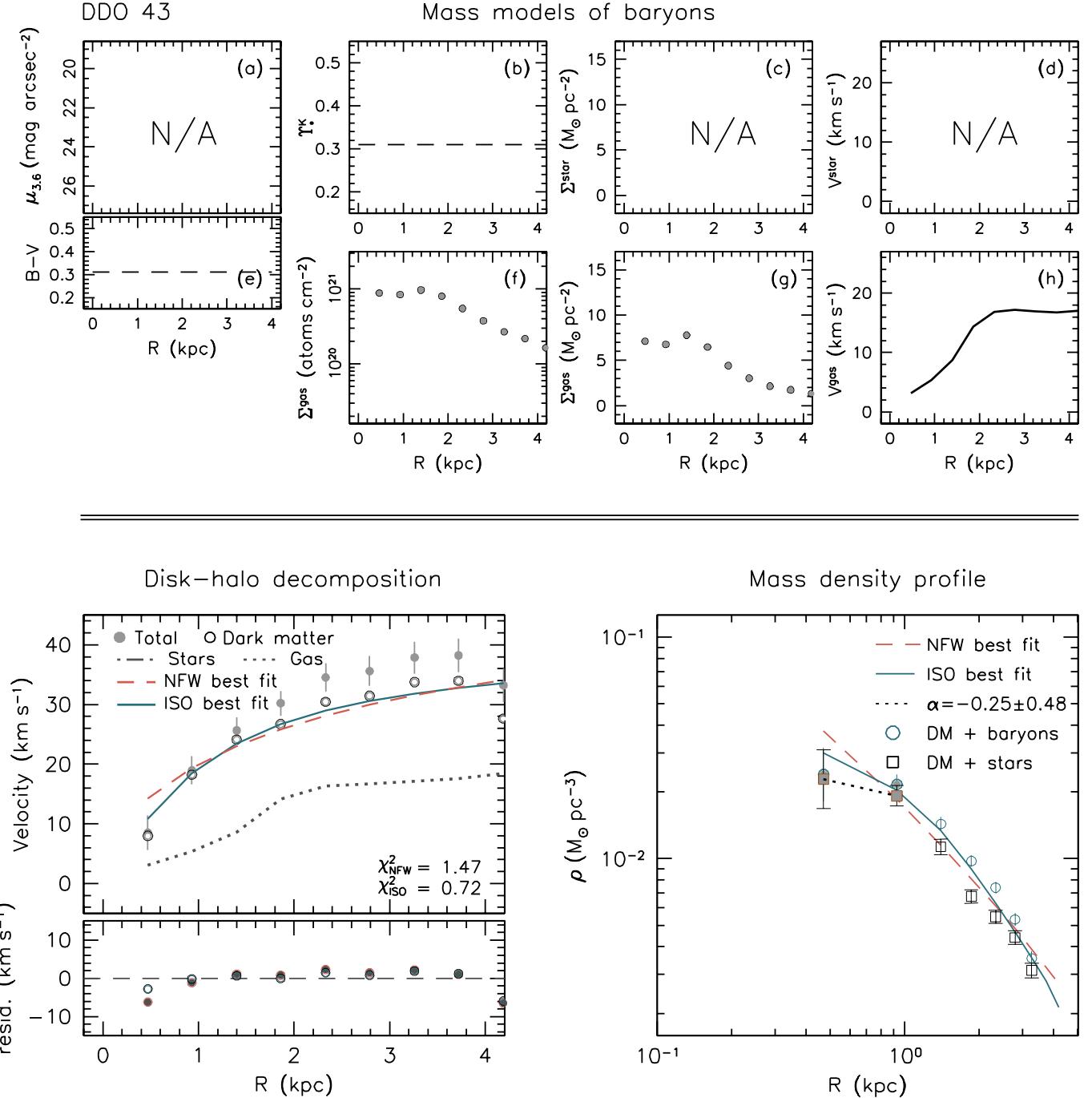


FIG. A.6.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 43. Please refer to the text in Sections 3 and 4 for full information.

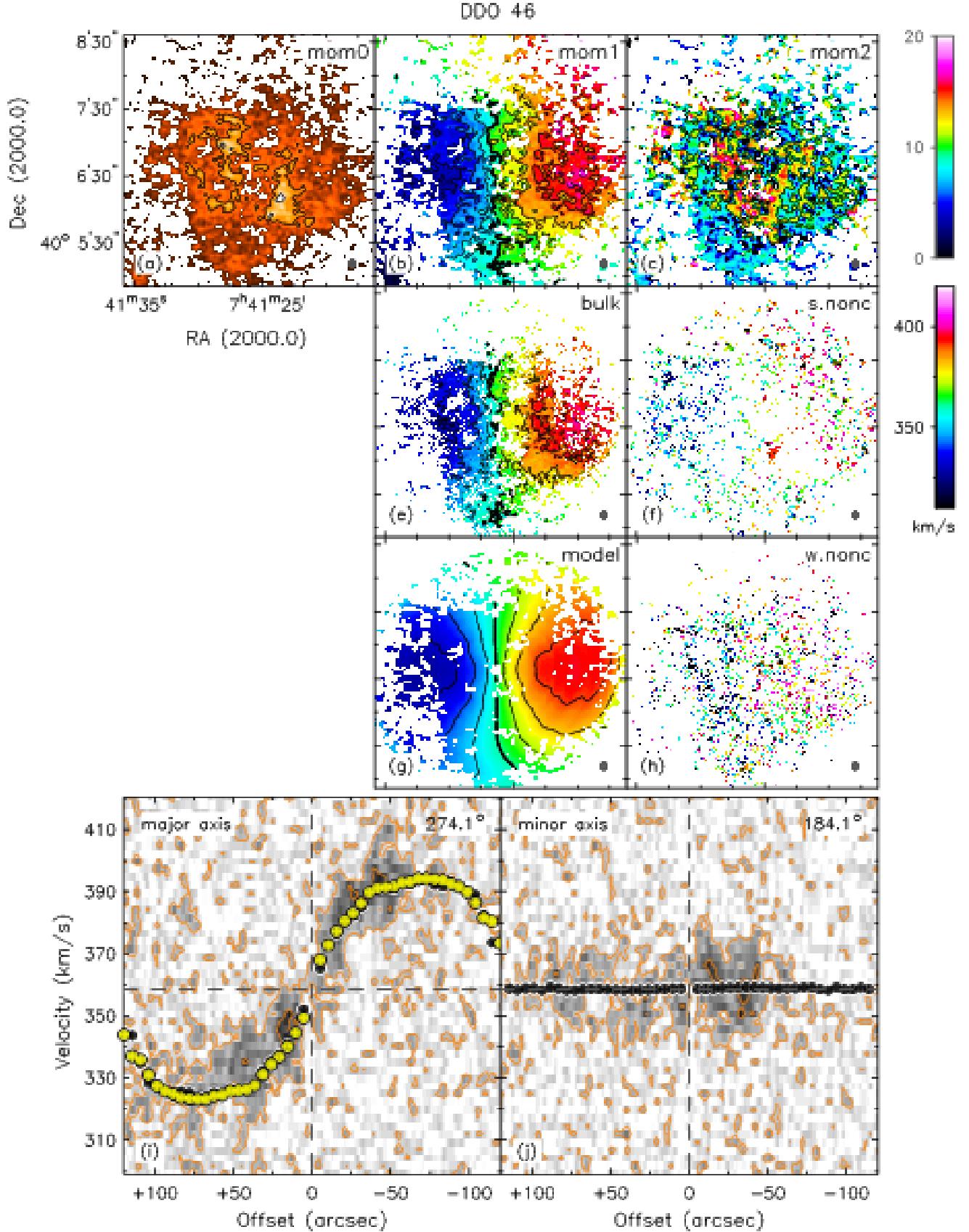


FIG. A.7.— HI data and *Spitzer* IRAC 3.6 μm image of DDO 46. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

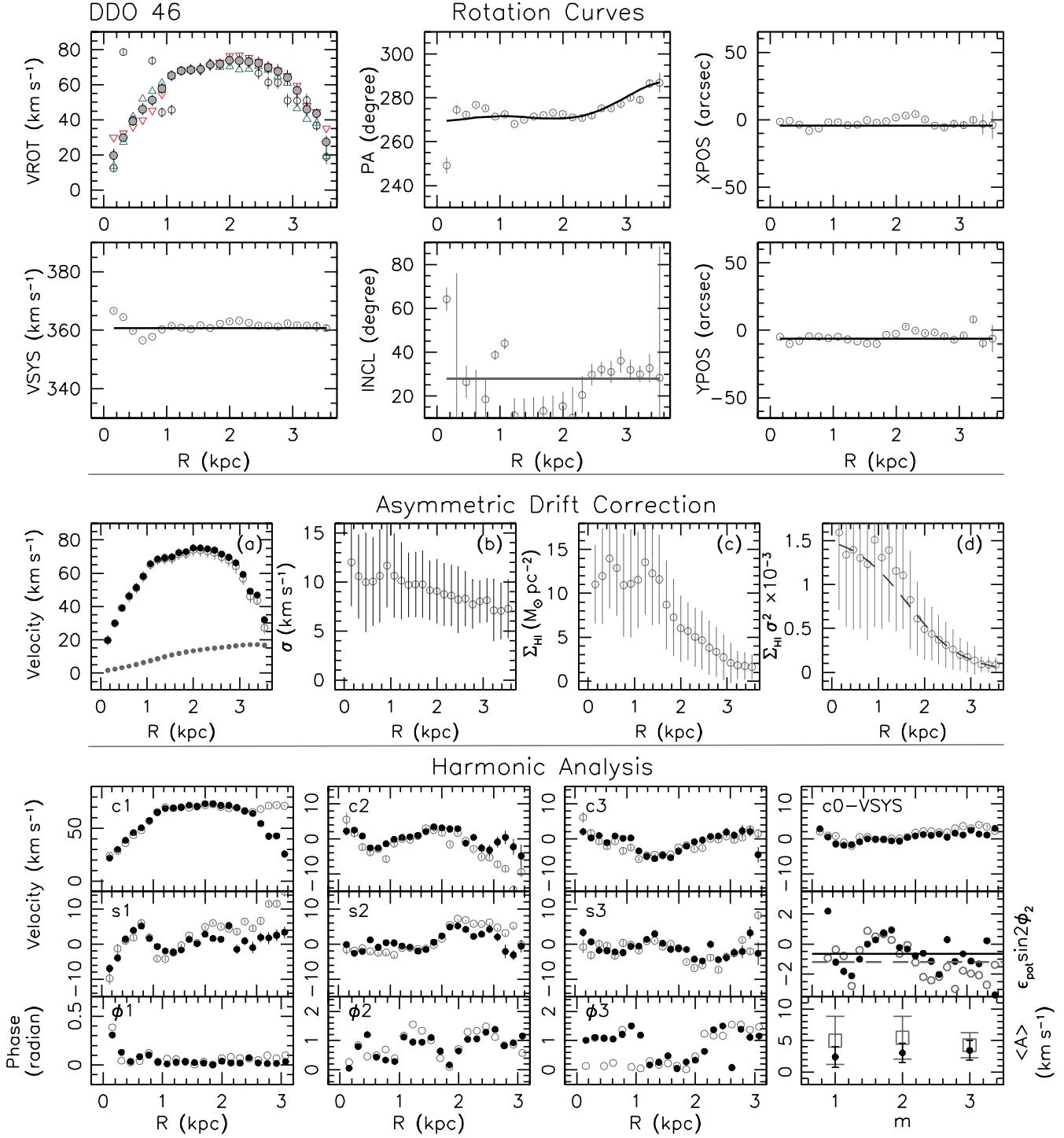


FIG. A.8.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 46. See Appendix section B for details.

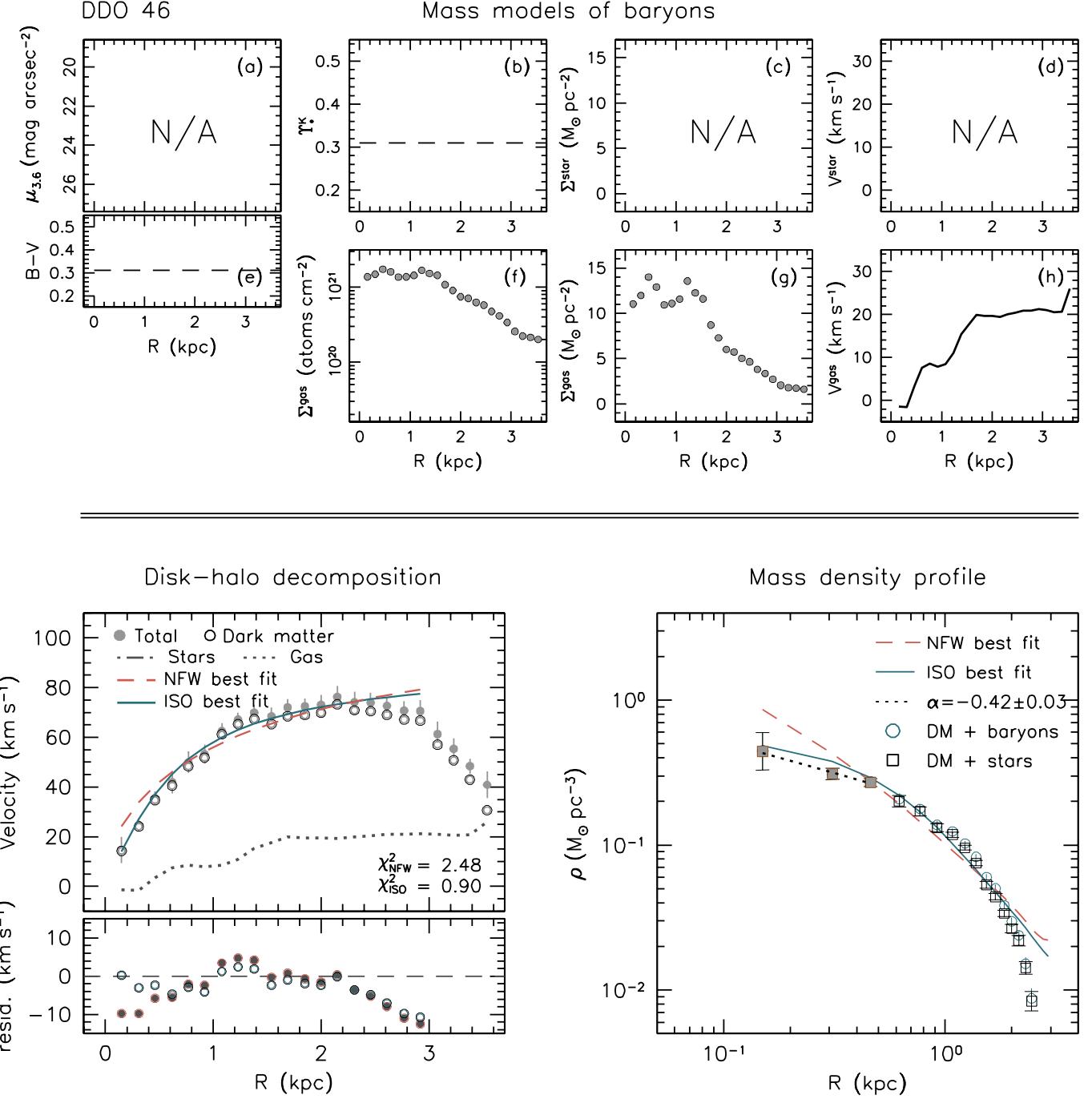


FIG. A.9.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 46. Please refer to the text in Sections 3 and 4 for full information.

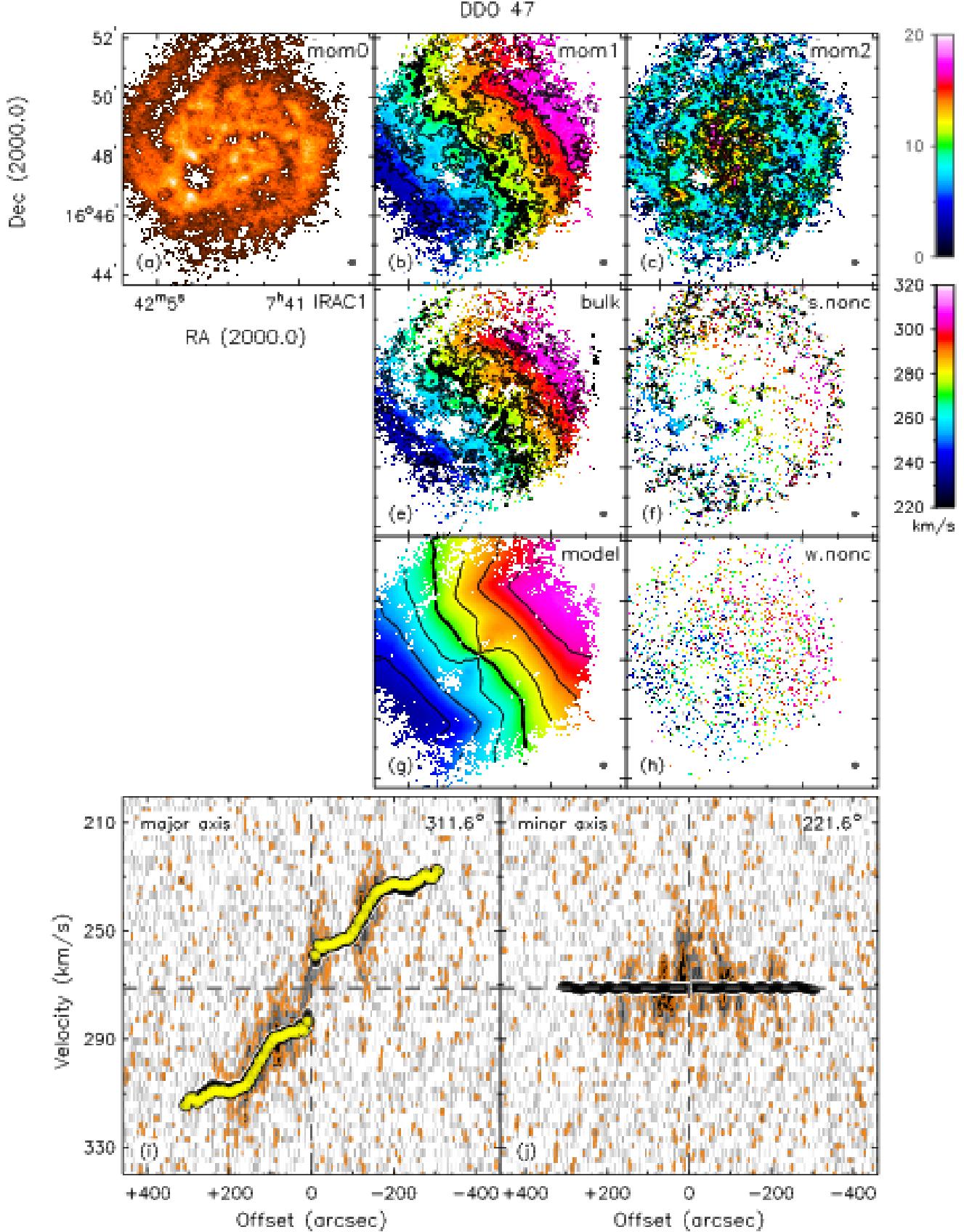


FIG. A.10.— HI data and *Spitzer* IRAC 3.6 μm image of DDO 47. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

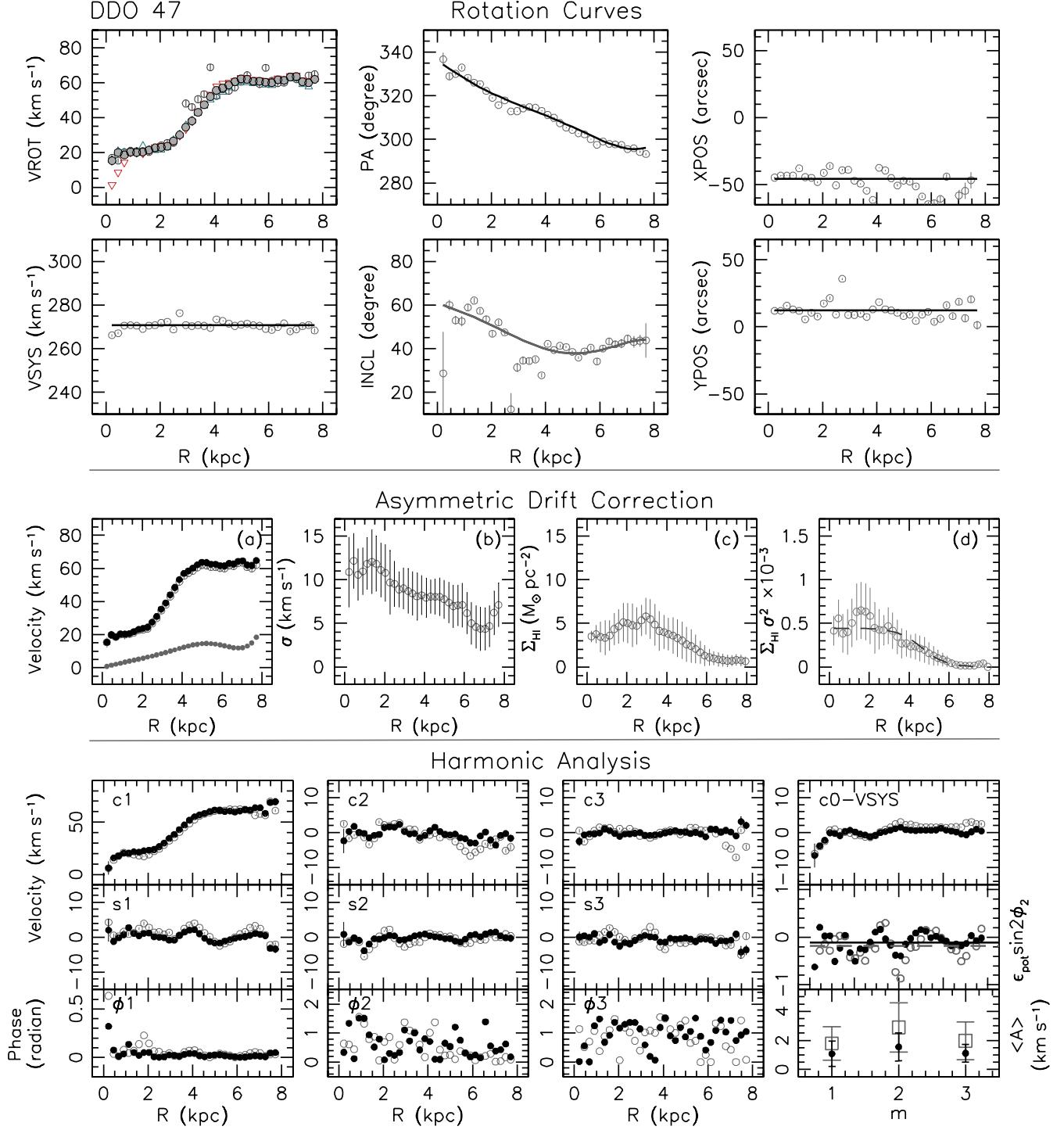
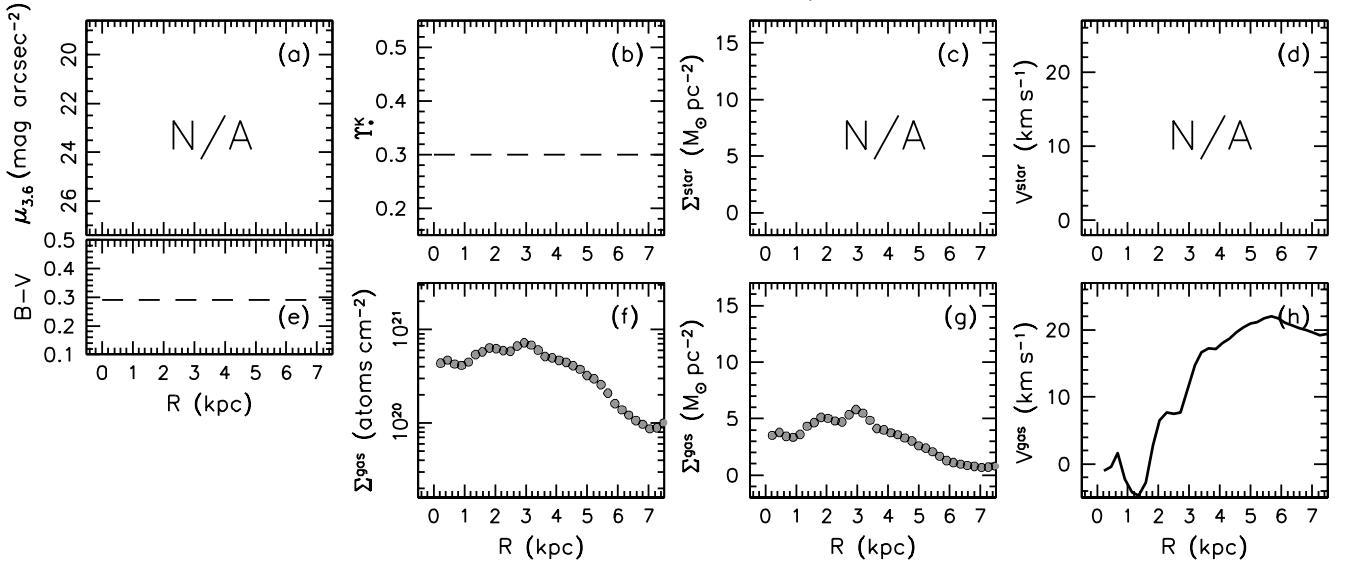


FIG. A.11.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 47. See Appendix section B for details.

DDO 47

Mass models of baryons



Disk-halo decomposition

Mass density profile

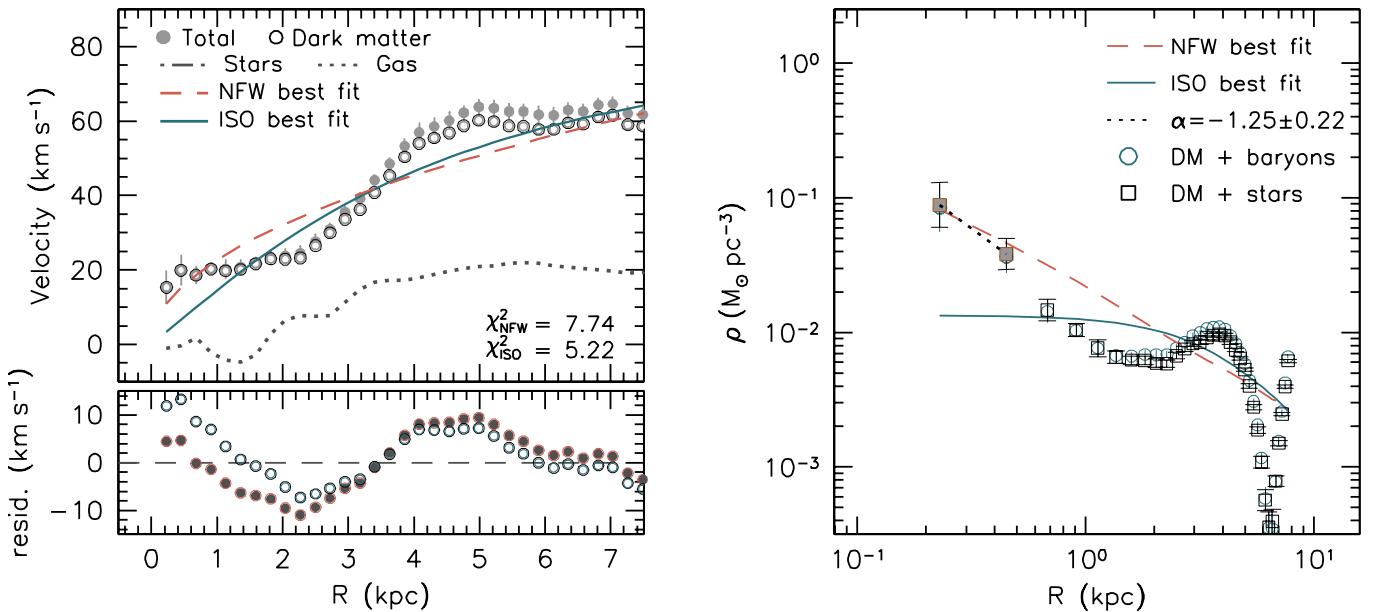


FIG. A.12.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 47. Please refer to the text in Sections 3 and 4 for full information.

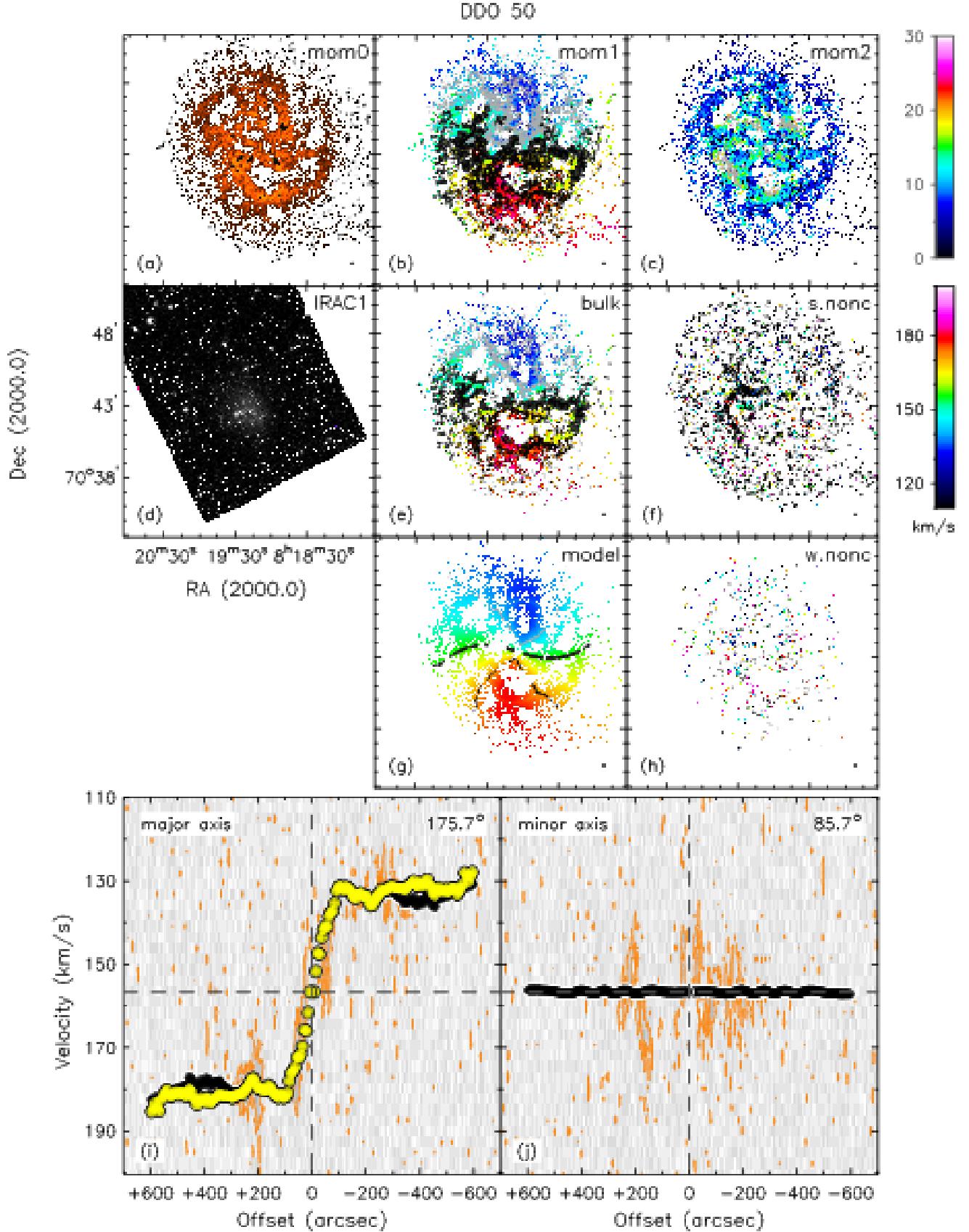


FIG. A.13.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 50. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 15 km s^{-1} . Velocity dispersion contours run from 0 to 30 km s^{-1} with a spacing of 20 km s^{-1} . See Appendix section A for details.

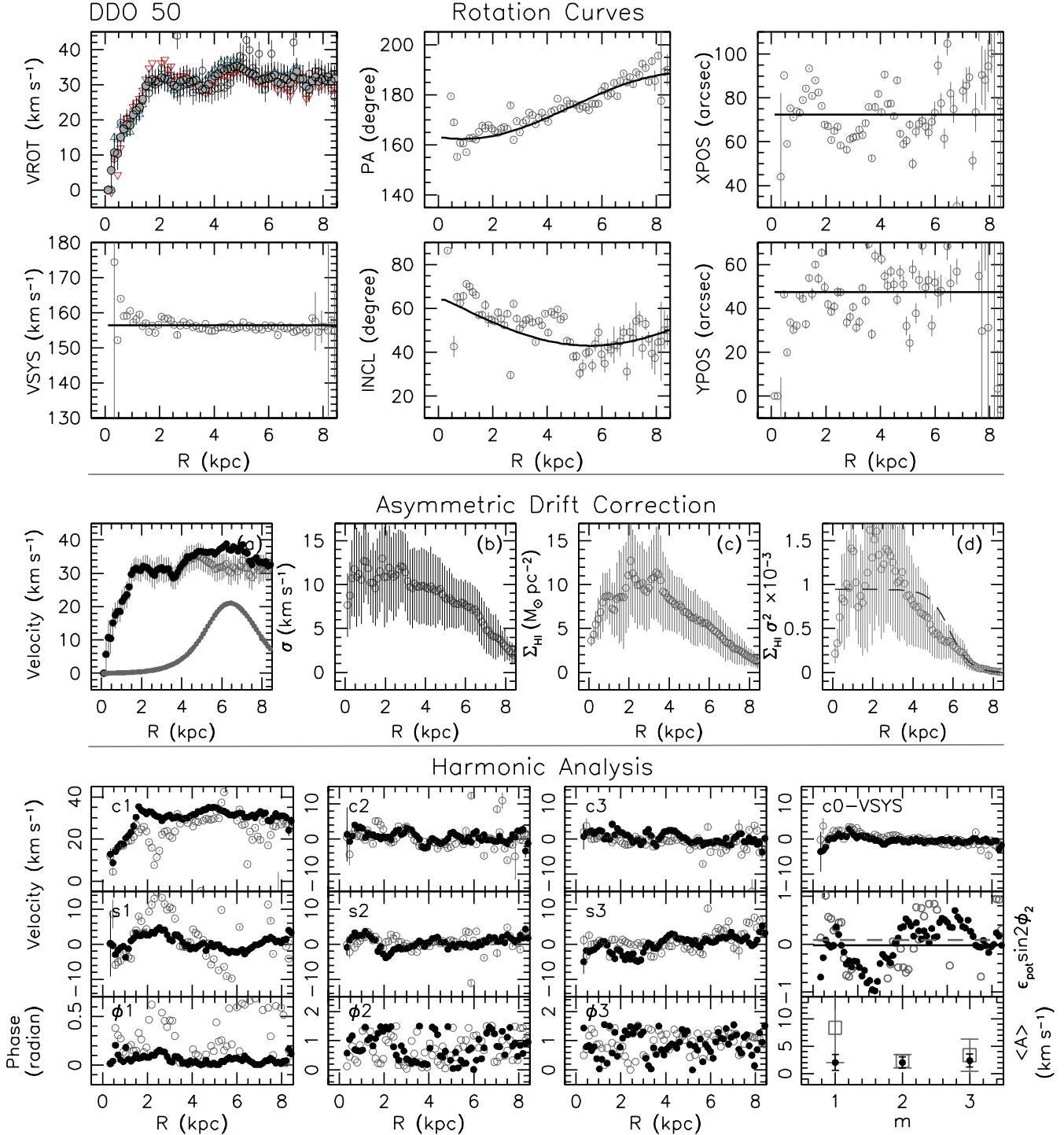


FIG. A.14.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 50. See Appendix section B for details.

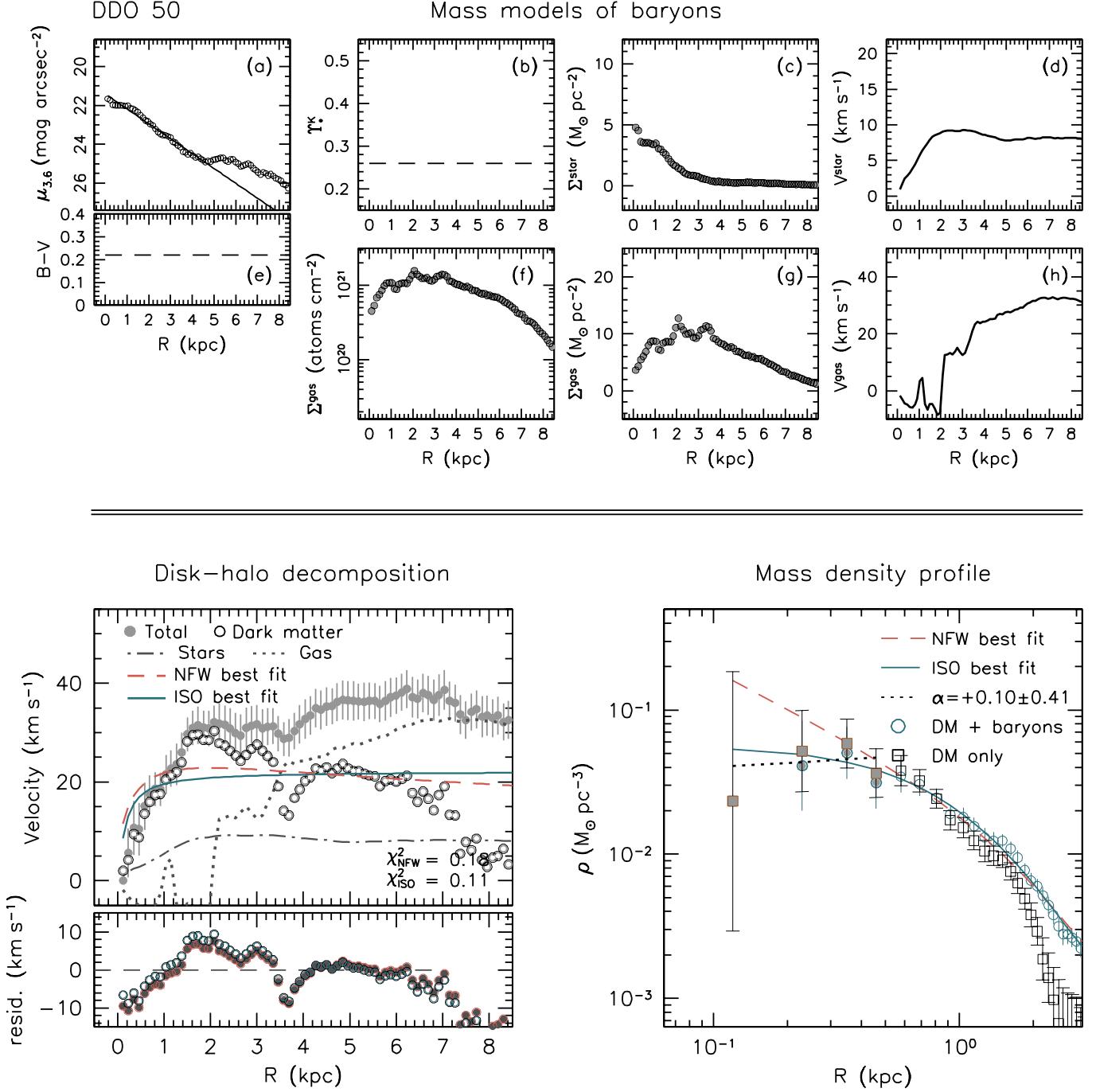


FIG. A.15.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 50. Please refer to the text in Sections 3 and 4 for full information.

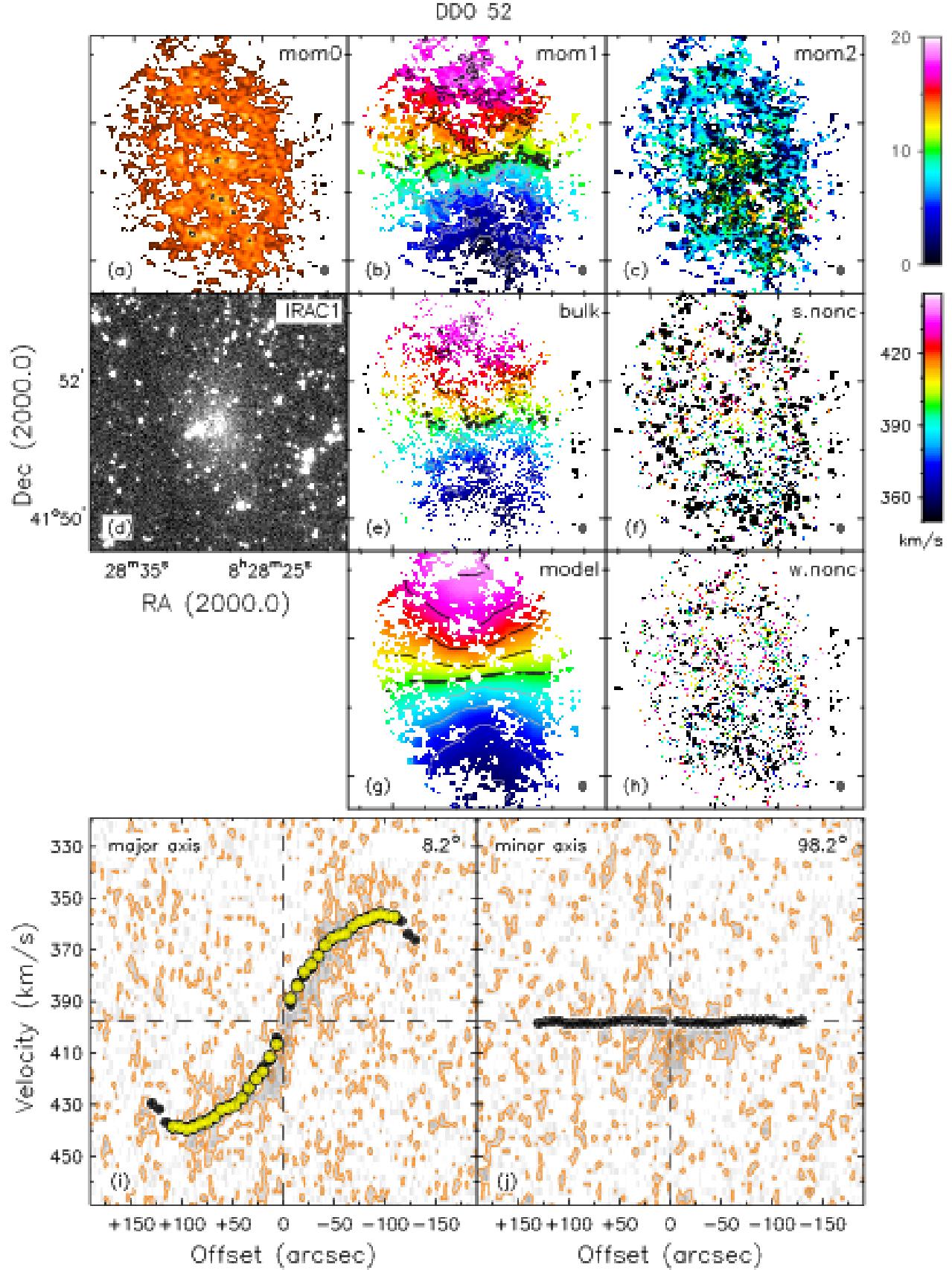


FIG. A.16.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 52. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

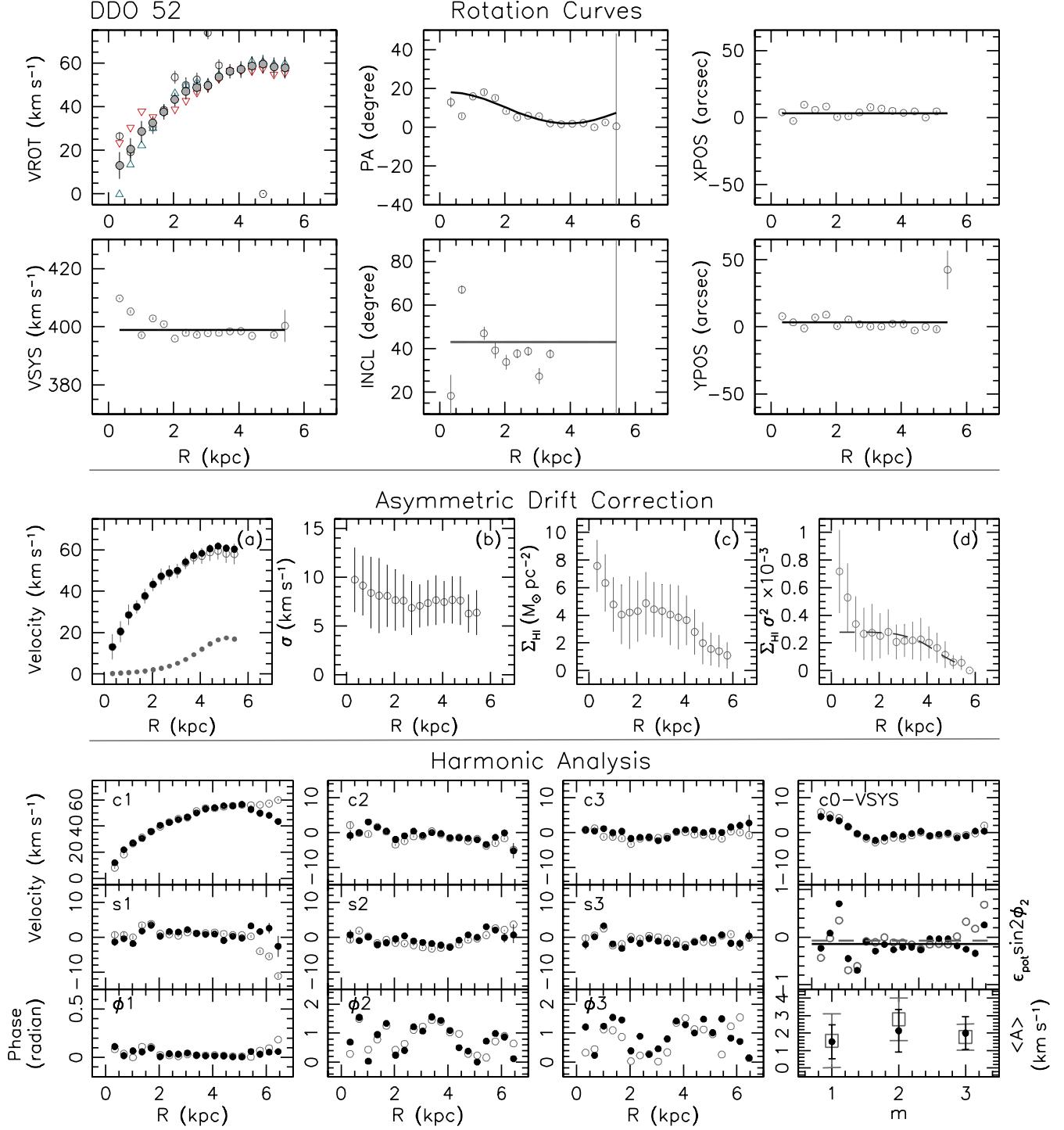


FIG. A.17.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 52. See Appendix section B for details.

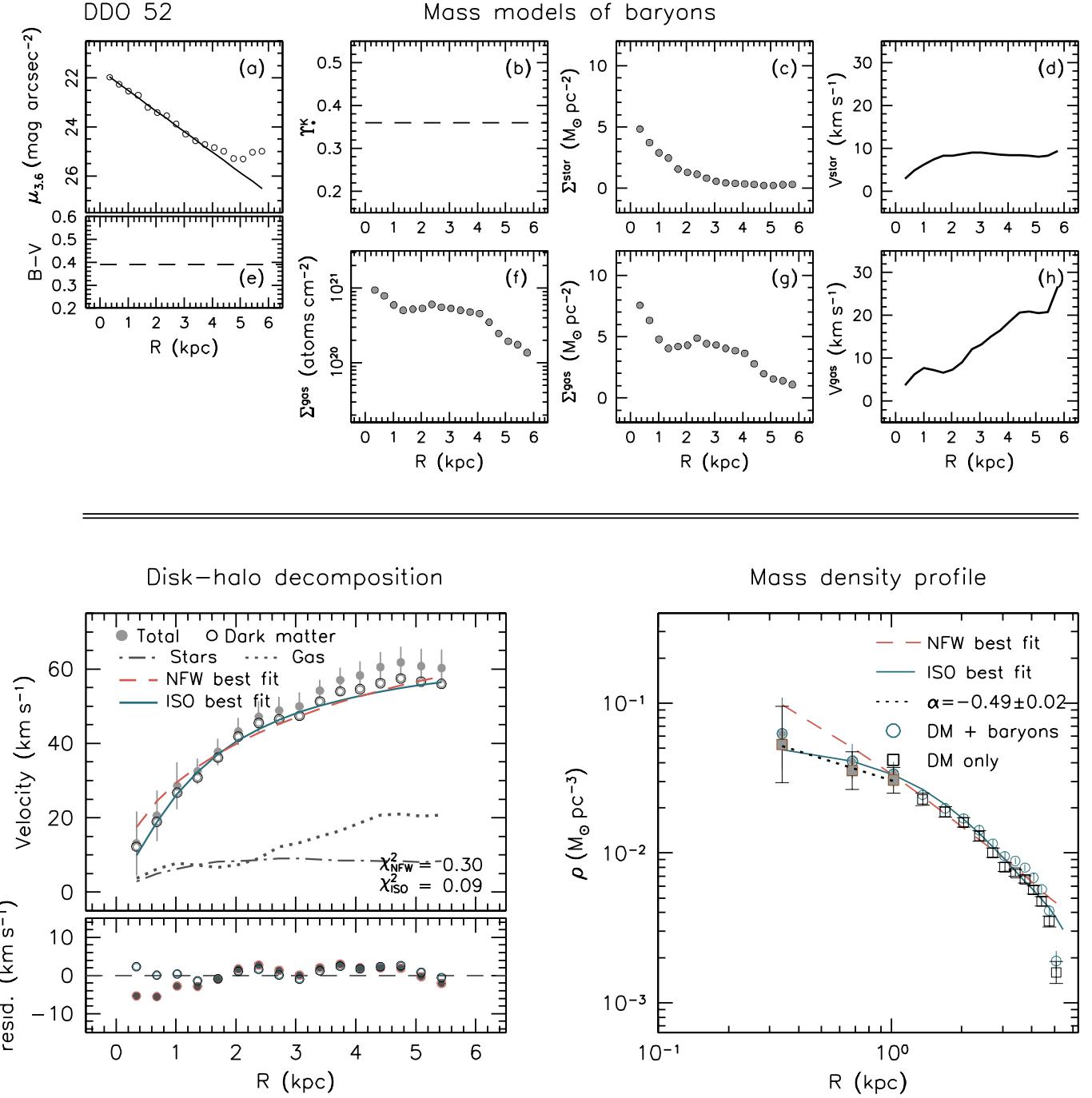


FIG. A.18.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 52. Please refer to the text in Sections 3 and 4 for full information.

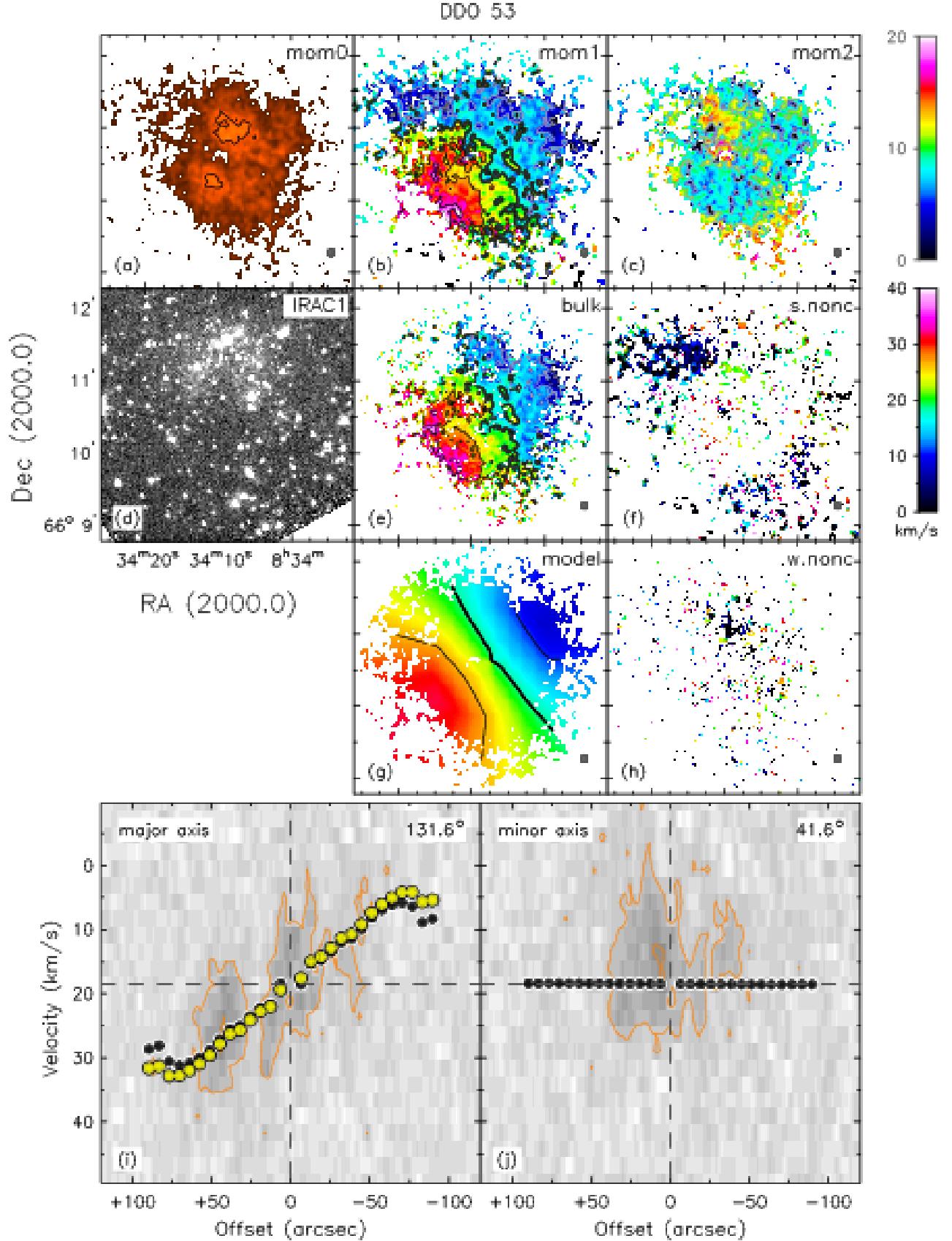


FIG. A.19.— HI data and *Spitzer* IRAC 3.6 μm image of DDO 53. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

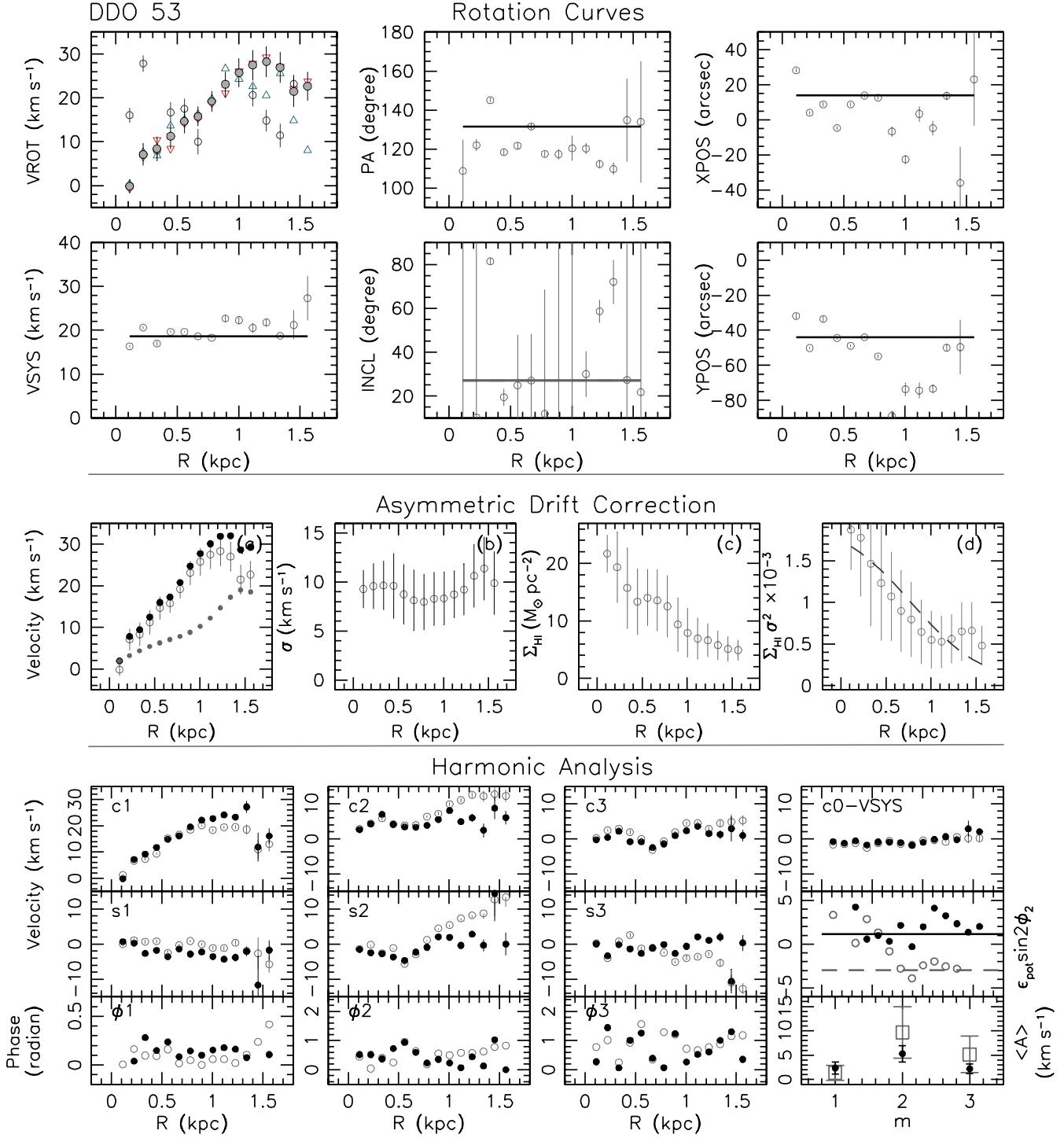


FIG. A.20.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 53. See Appendix section B for details.

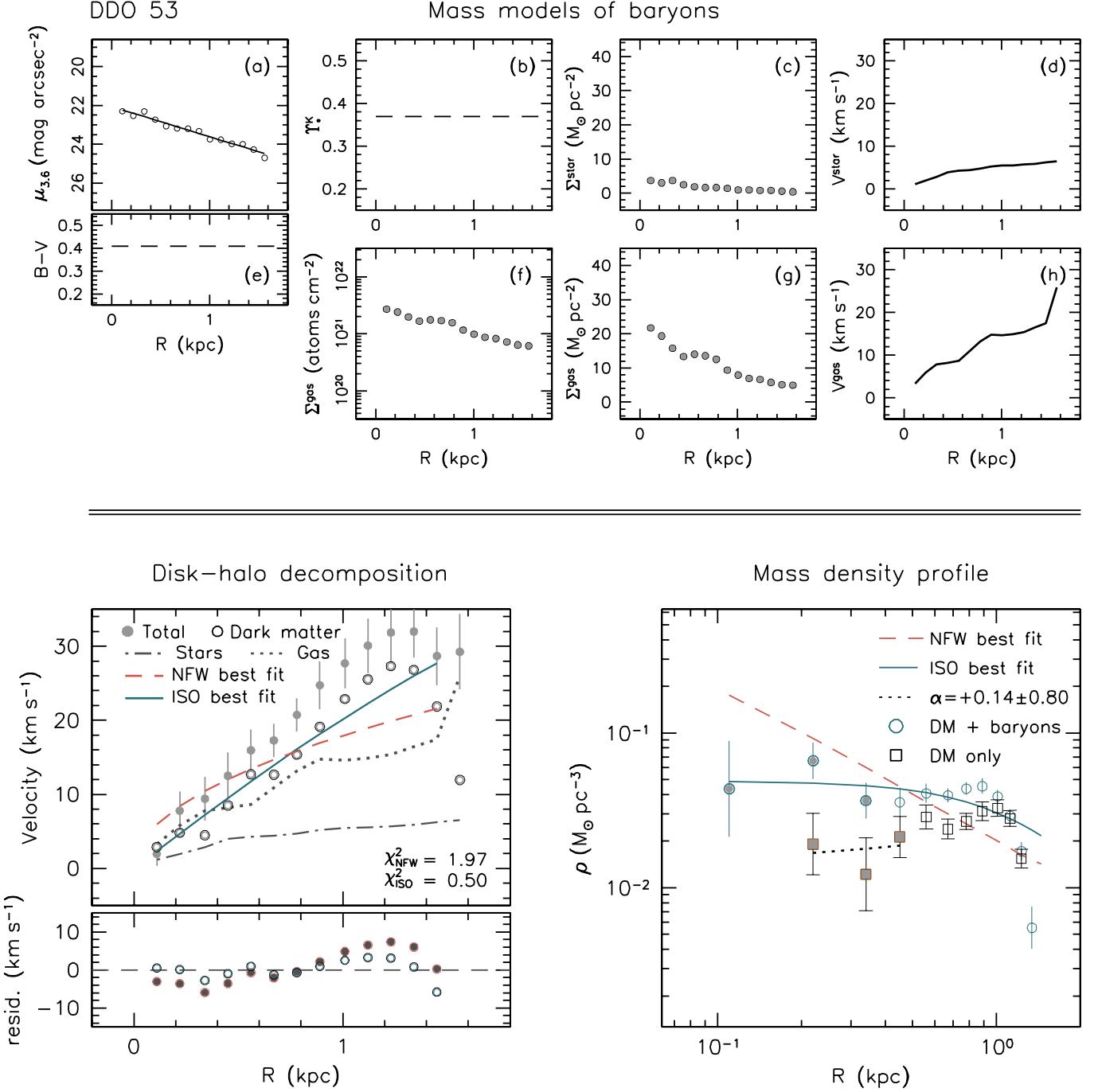


FIG. A.21.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 53. Please refer to the text in Sections 3 and 4 for full information.

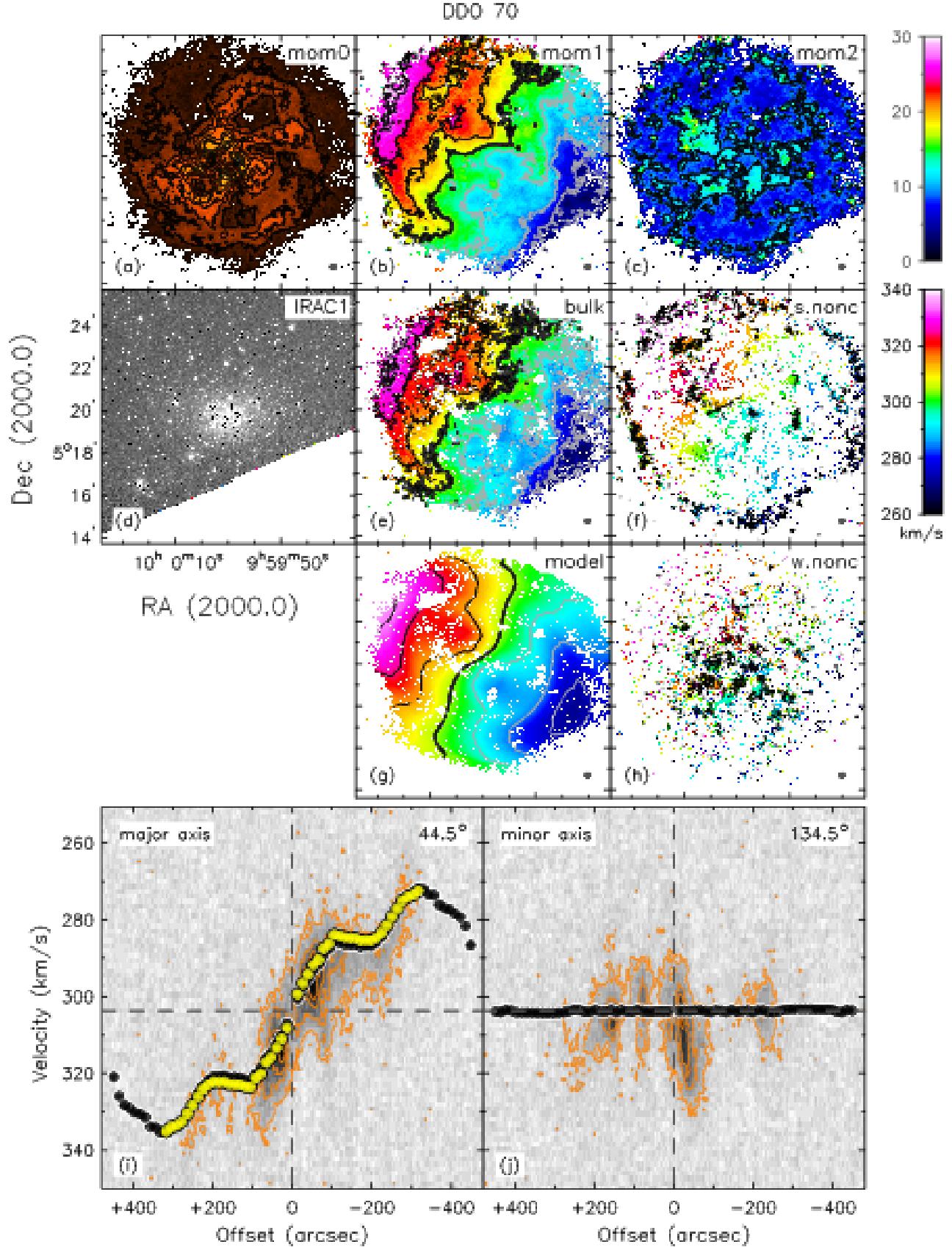


FIG. A.22.— HI data and *Spitzer* IRAC 3.6 μm image of DDO 70. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 30 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

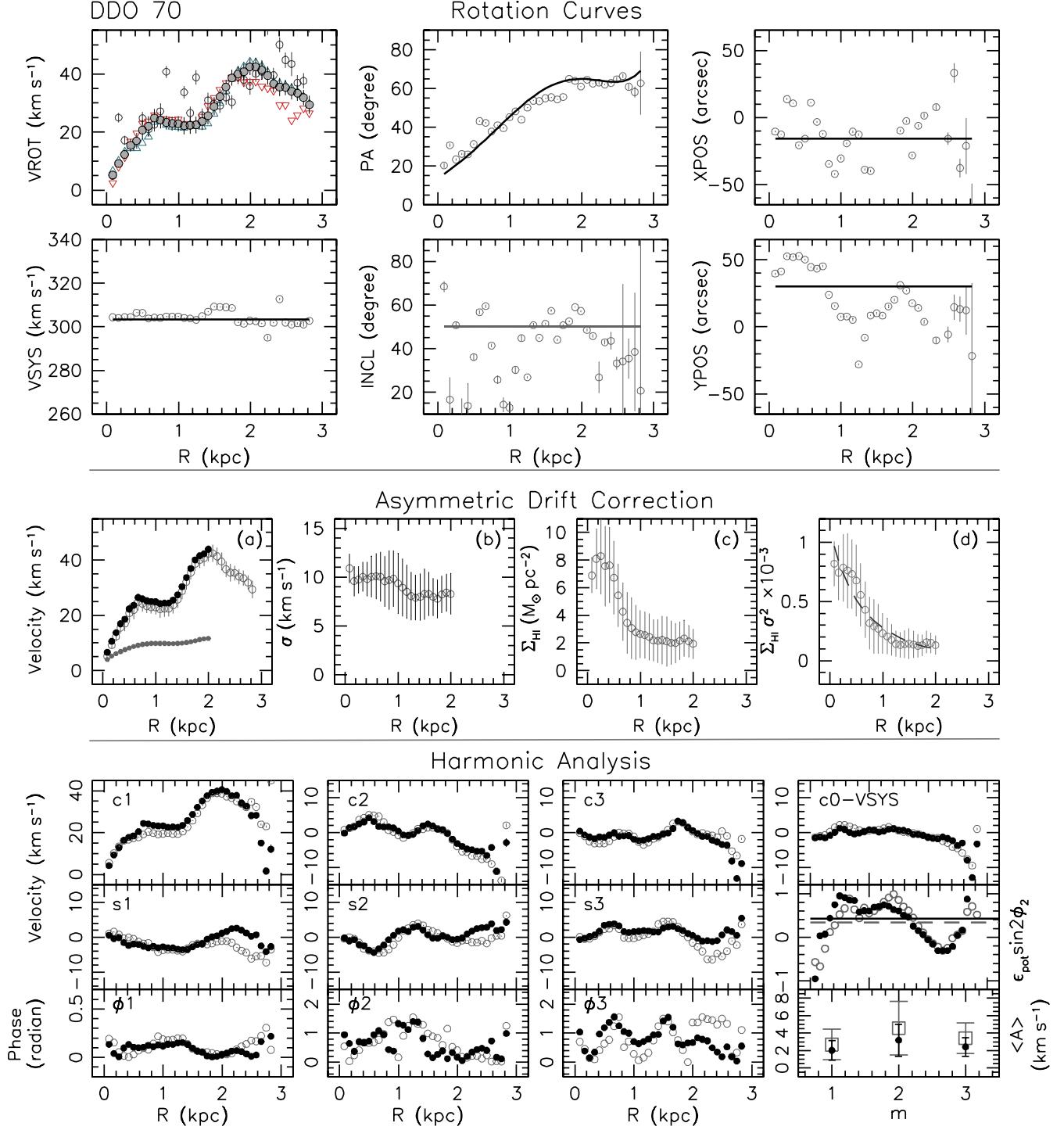


FIG. A.23.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 70. See Appendix section B for details.

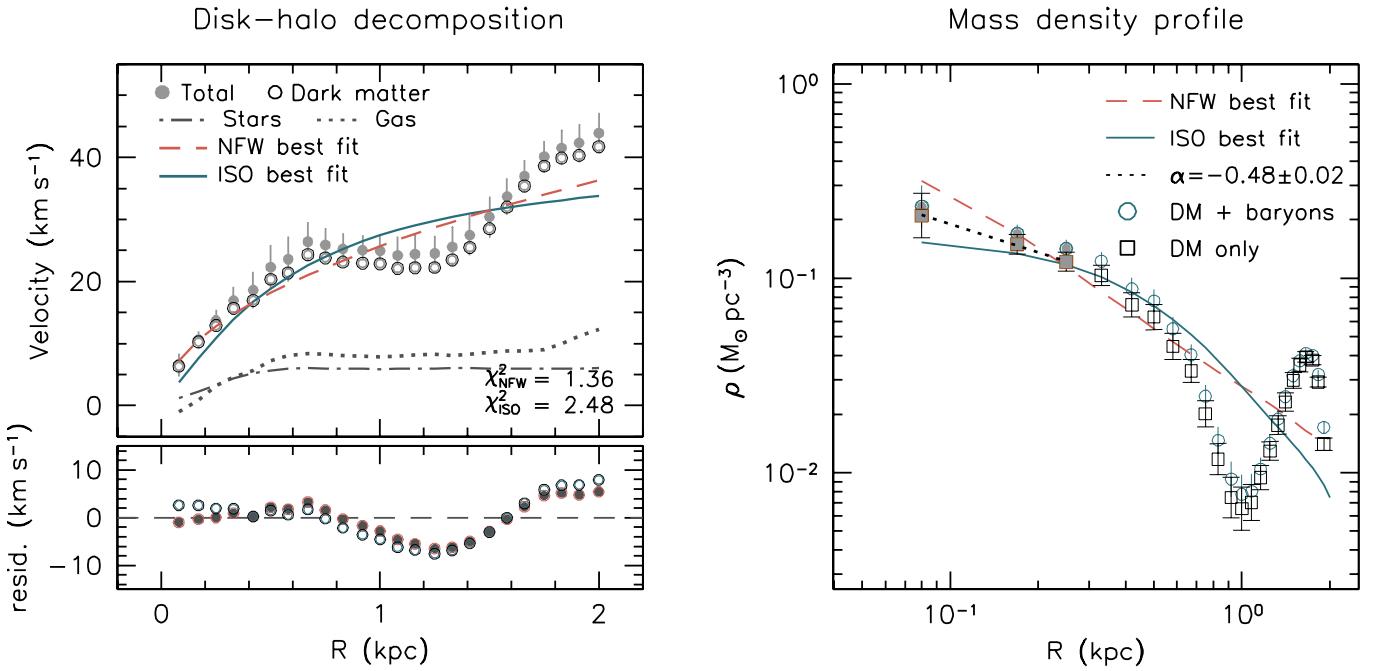
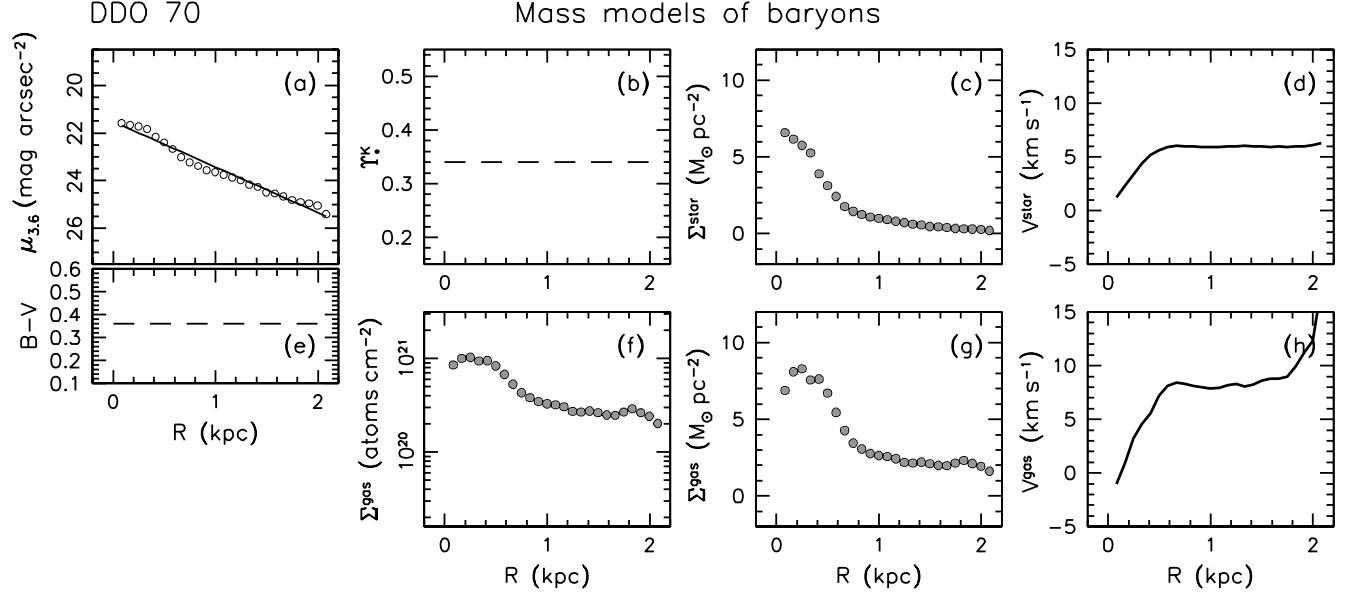


FIG. A.24.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 70. Please refer to the text in Sections 3 and 4 for full information.

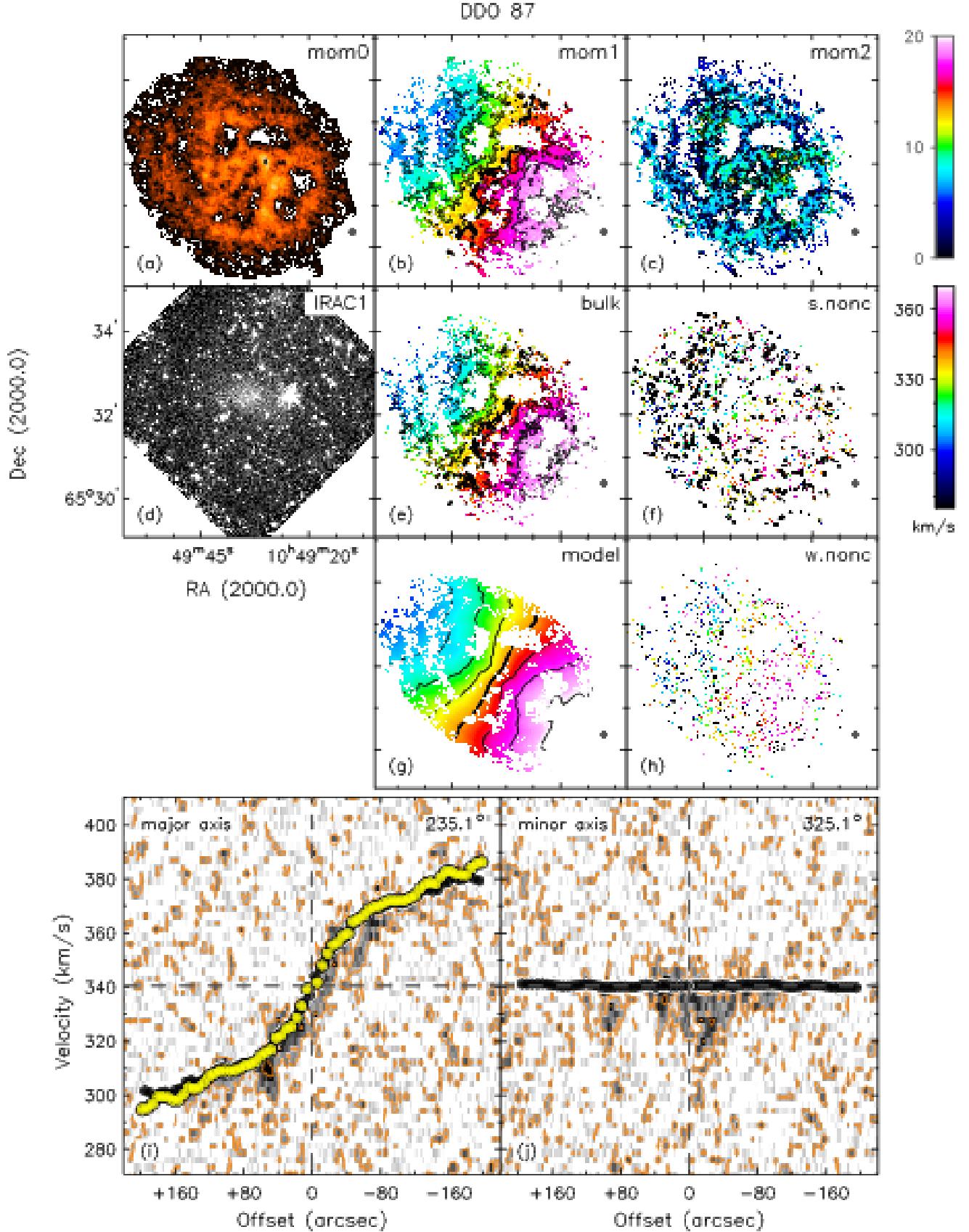


FIG. A.25.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 87. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 5 km s $^{-1}$. See Appendix section A for details.

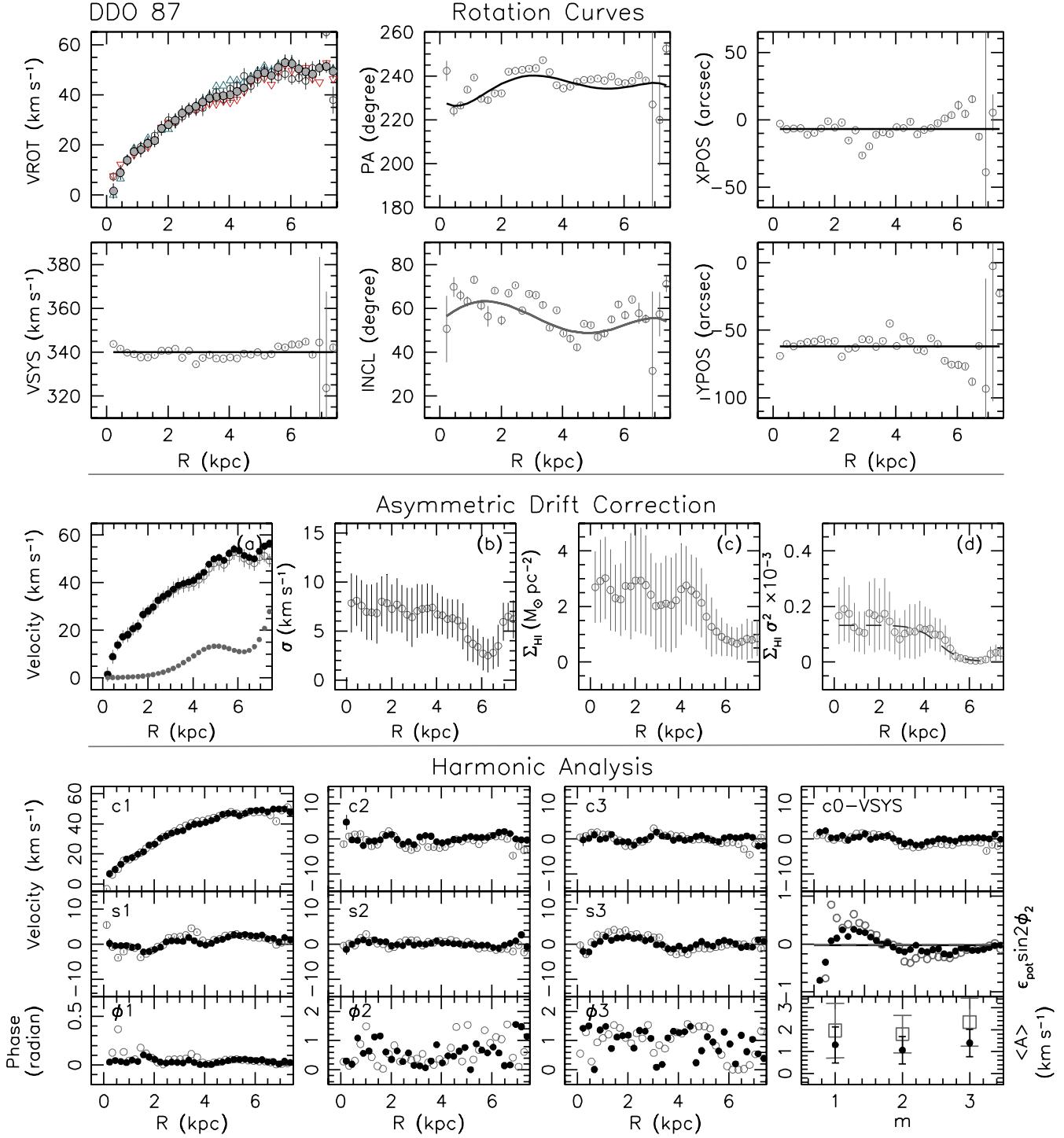


FIG. A.26.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 87. See Appendix section B for details.

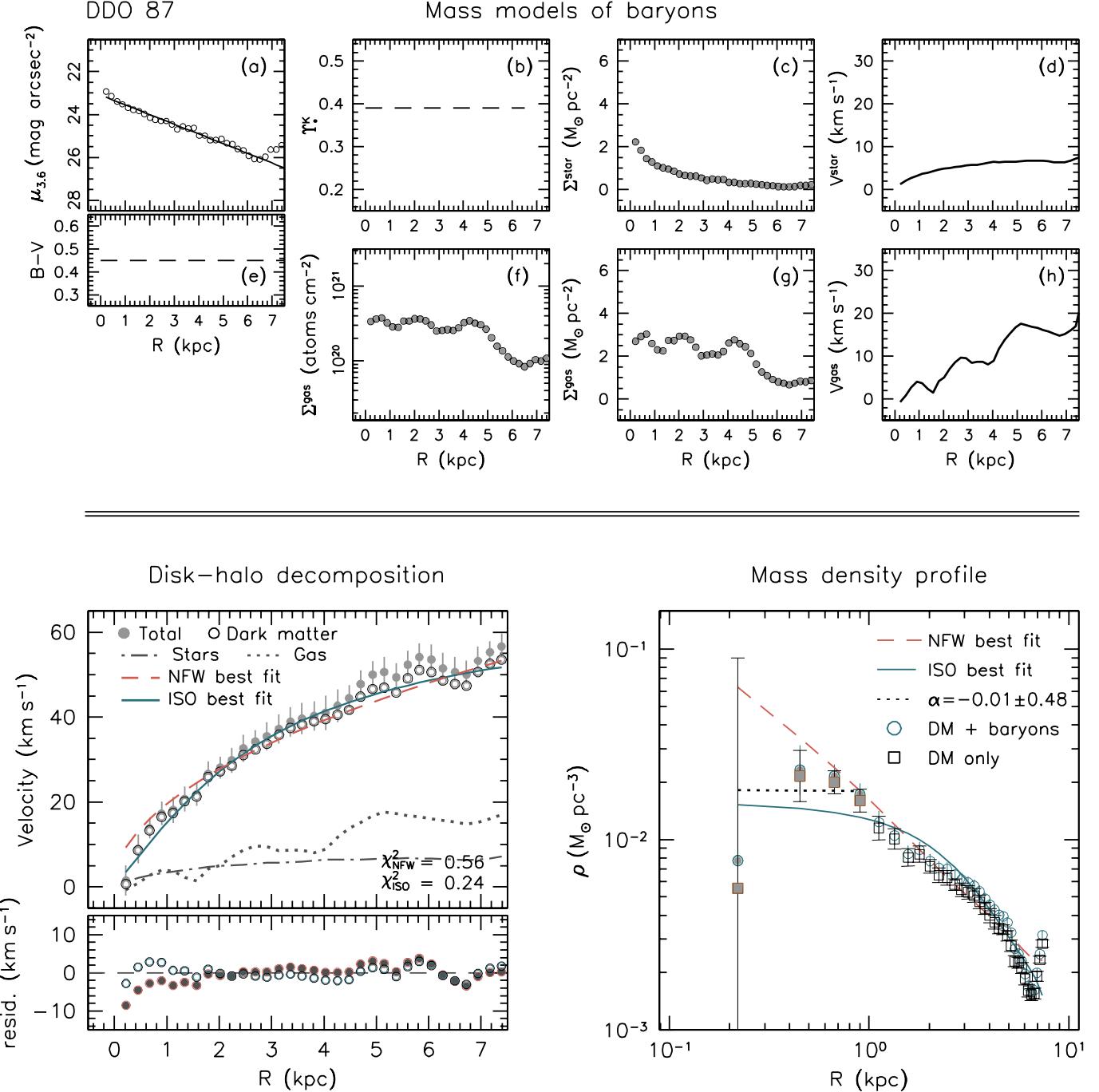


FIG. A.27.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 87. Please refer to the text in Sections 3 and 4 for full information.

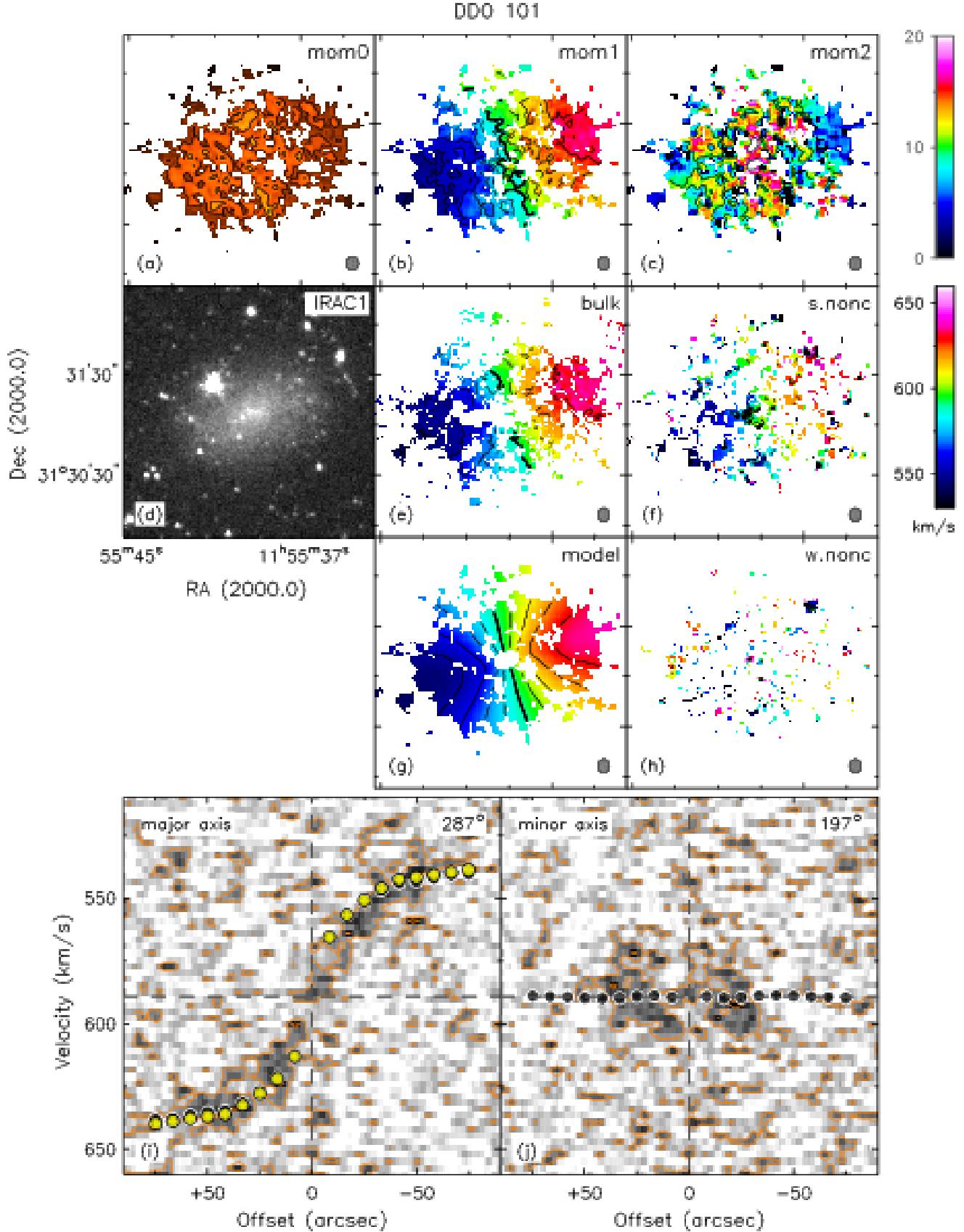


FIG. A.28.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 101. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

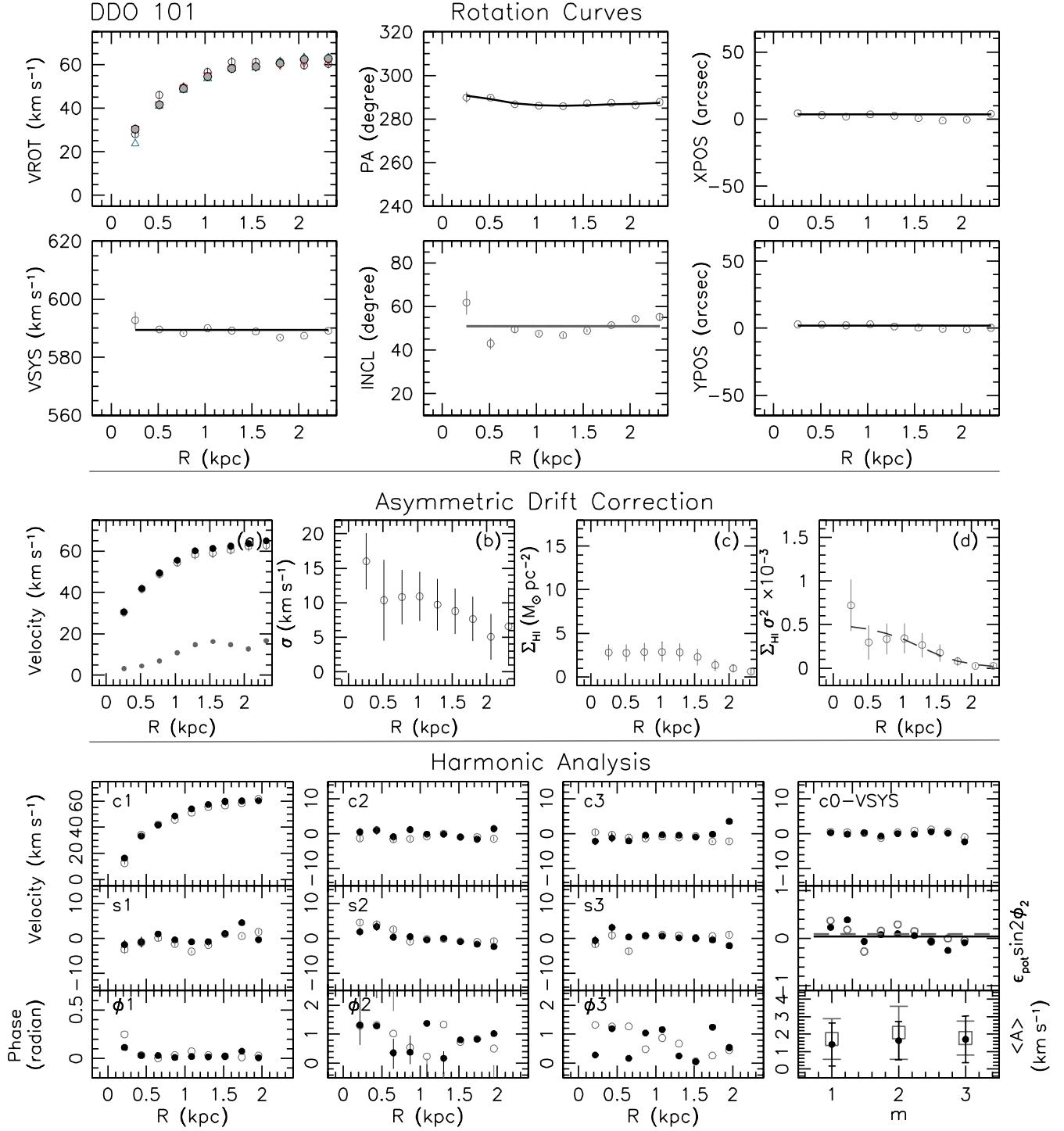


FIG. A.29.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 101. See Appendix section B for details.

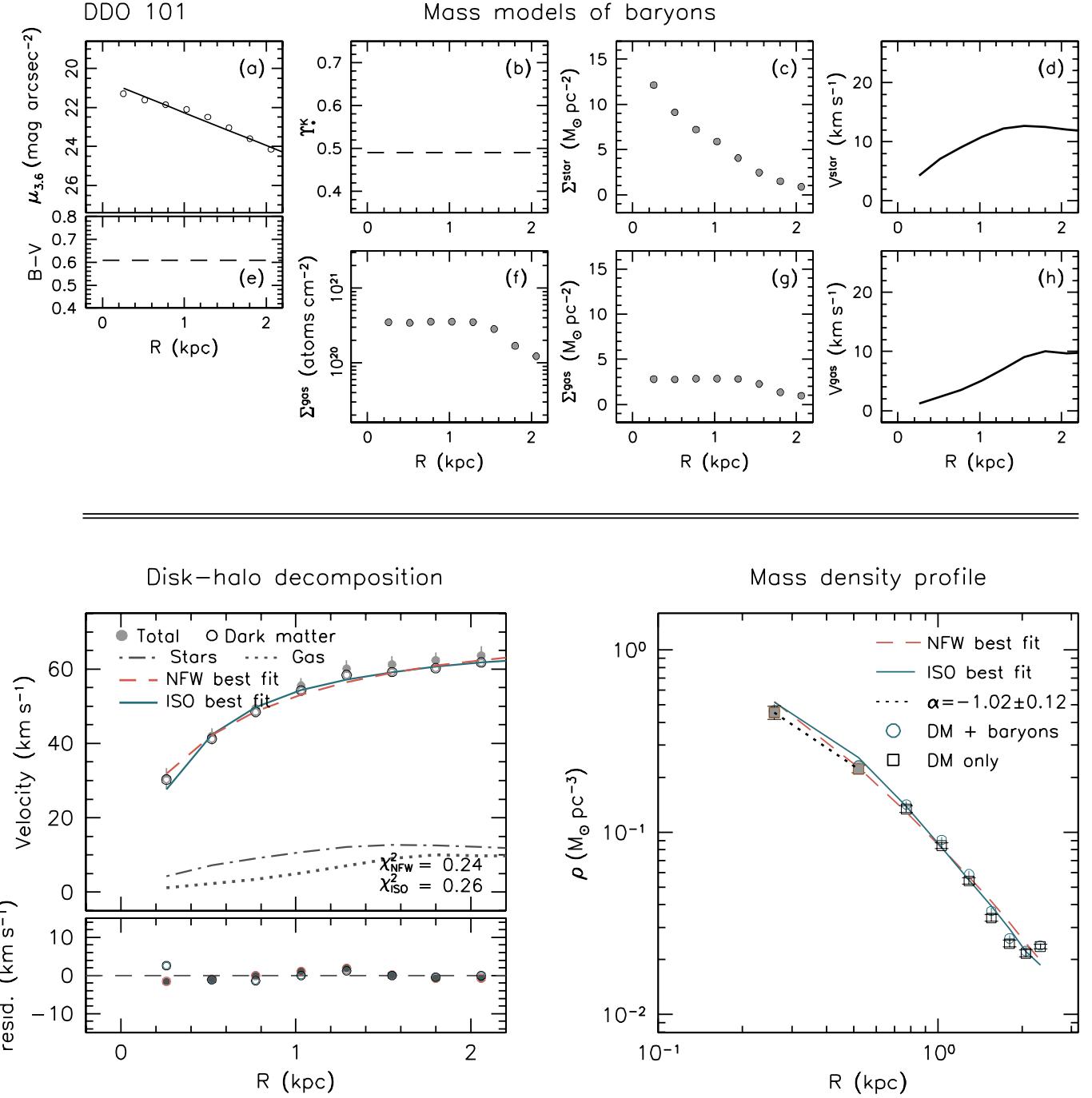


FIG. A.30.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 101. Please refer to the text in Sections 3 and 4 for full information.

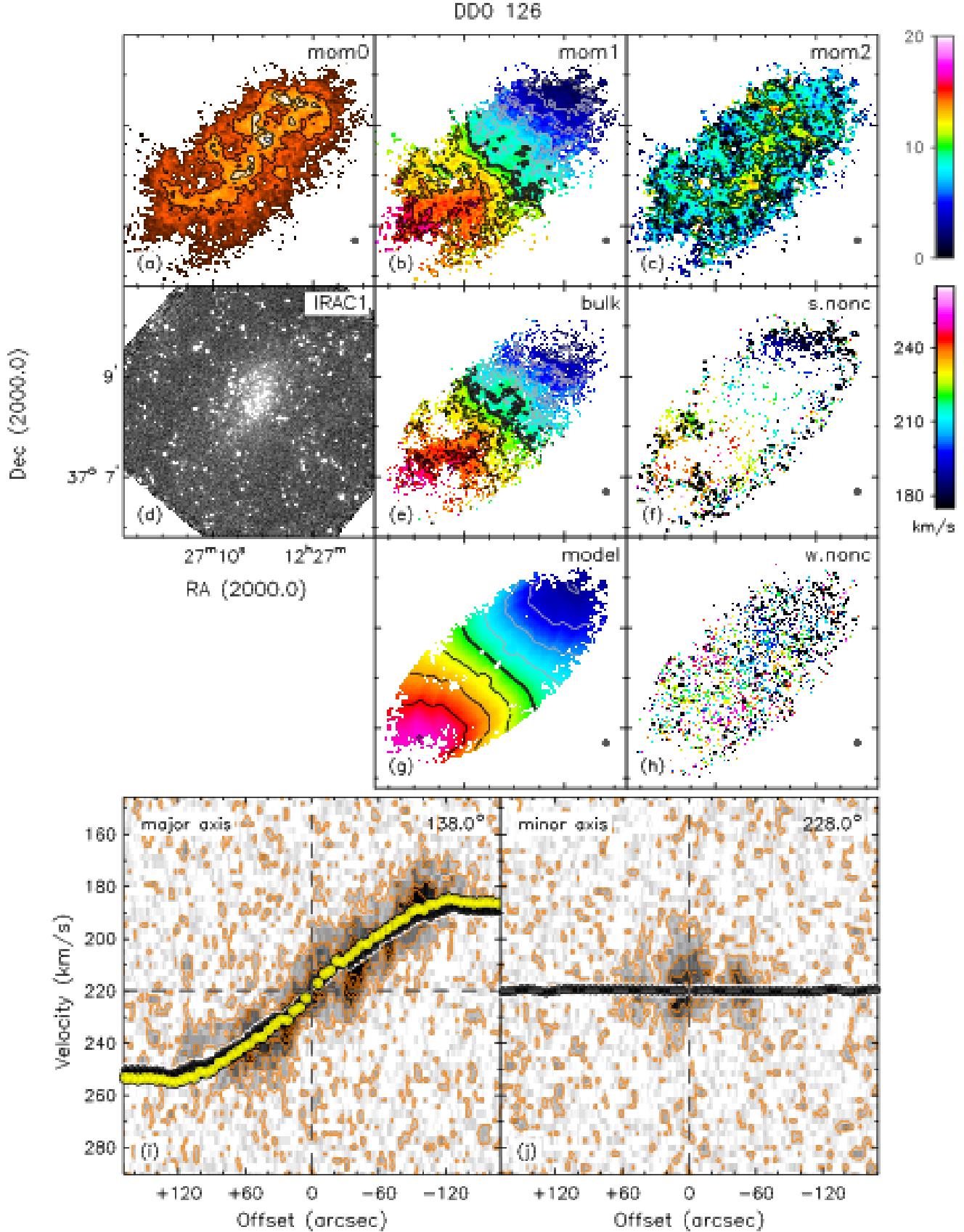


FIG. A.31.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 126. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 5 km s $^{-1}$. See Appendix section A for details.

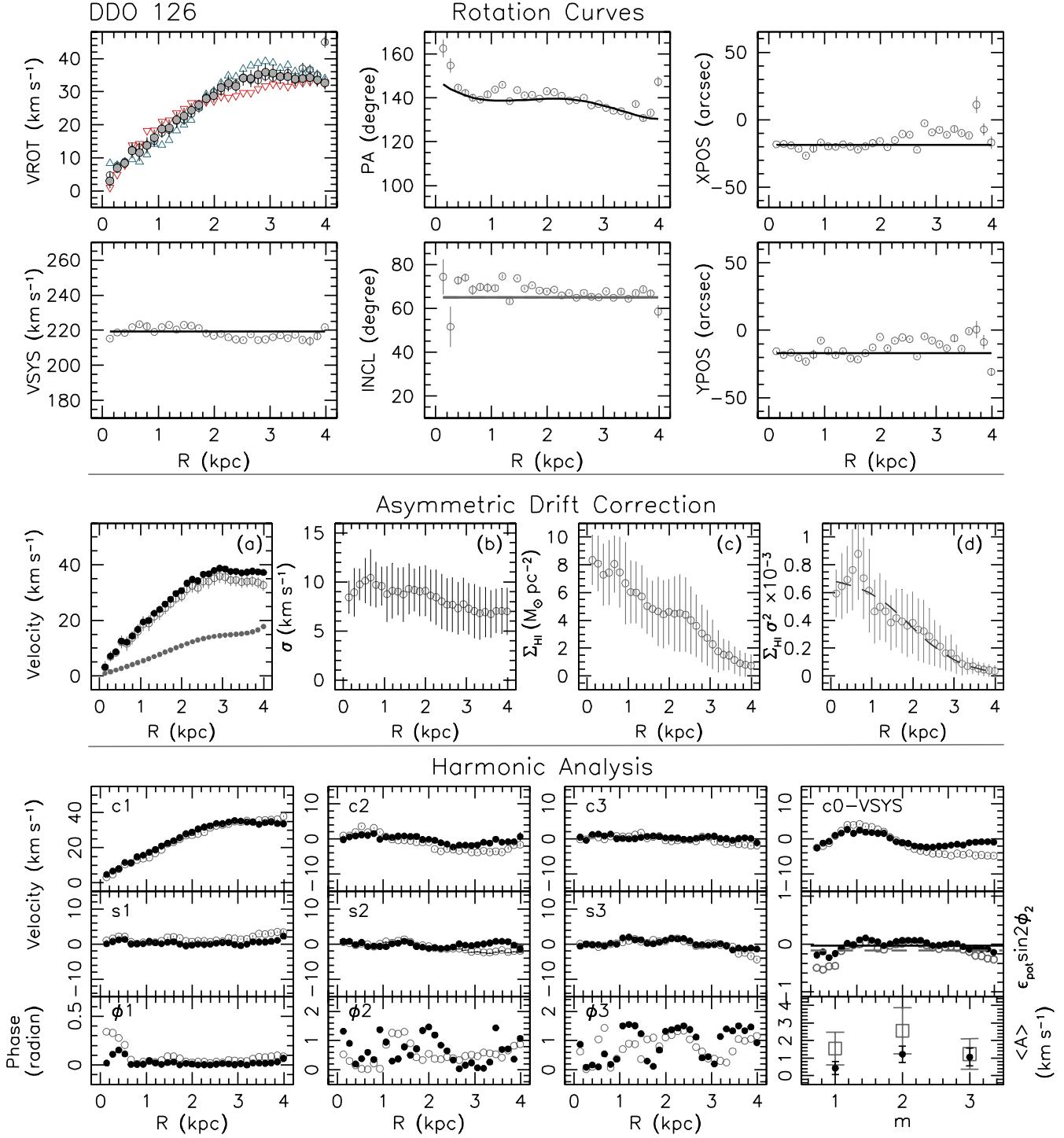


FIG. A.32.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 126. See Appendix section B for details.

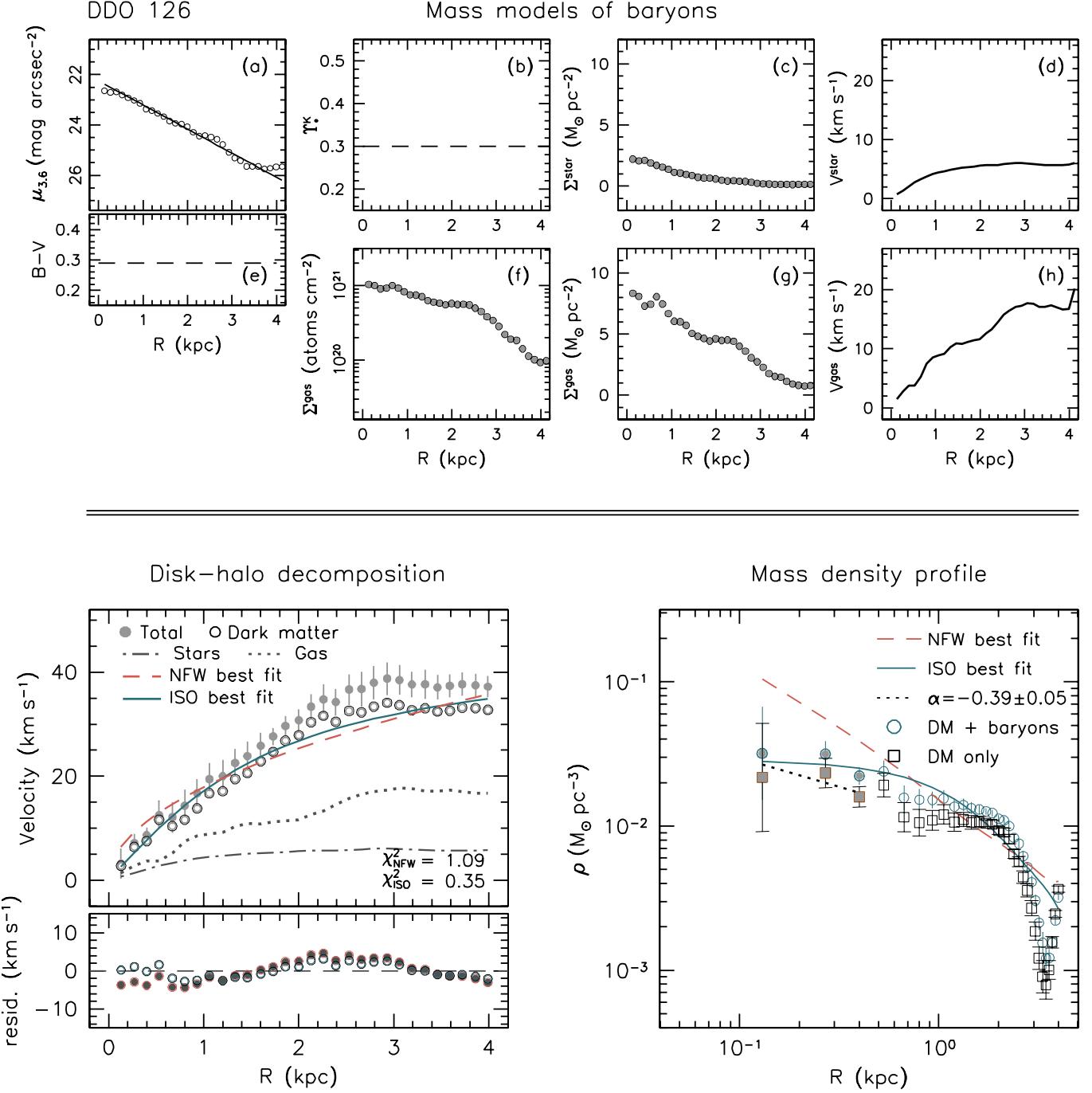


FIG. A.33.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 126. Please refer to the text in Sections 3 and 4 for full information.

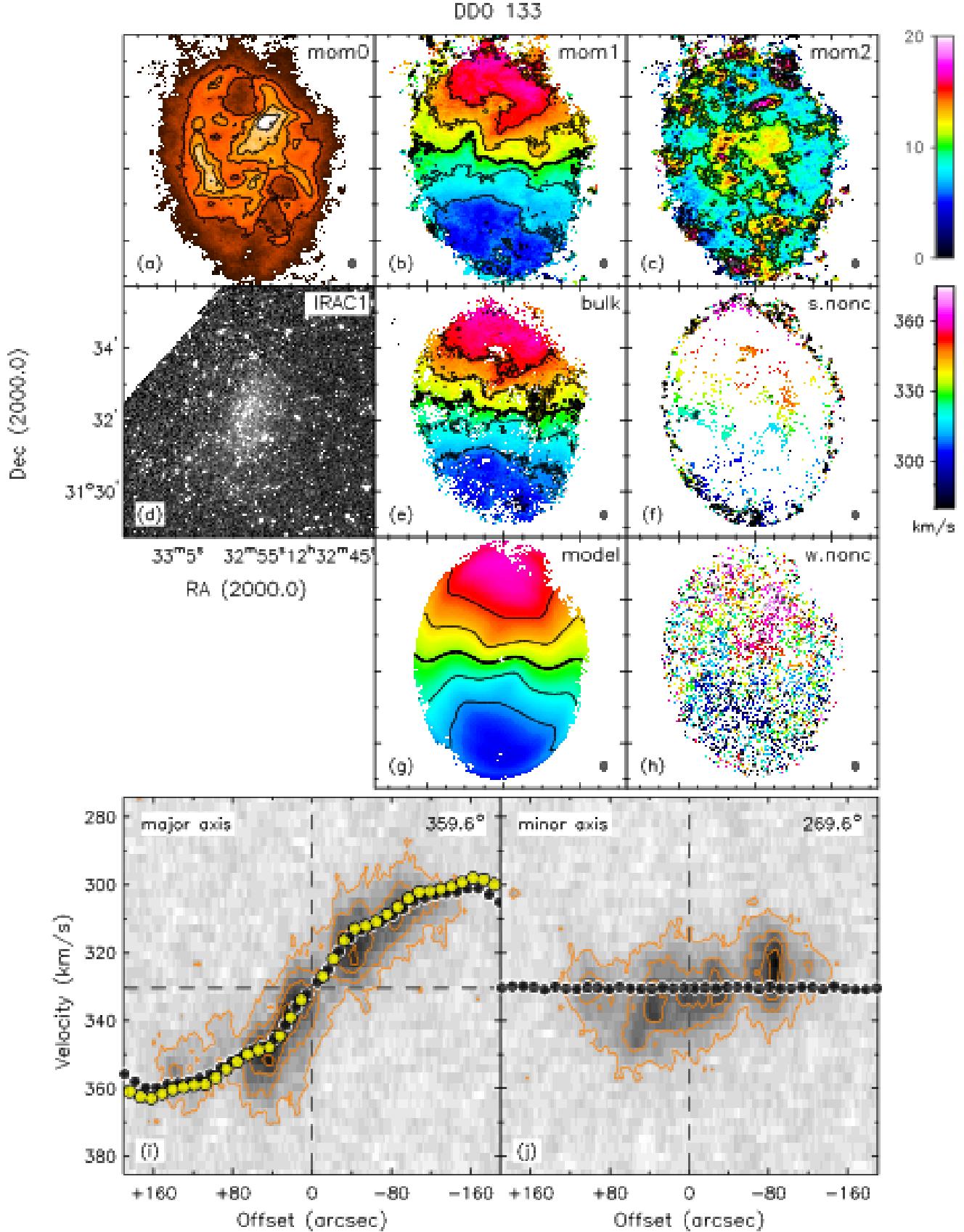


FIG. A.34.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 133. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 5 km s $^{-1}$. See Appendix section A for details.

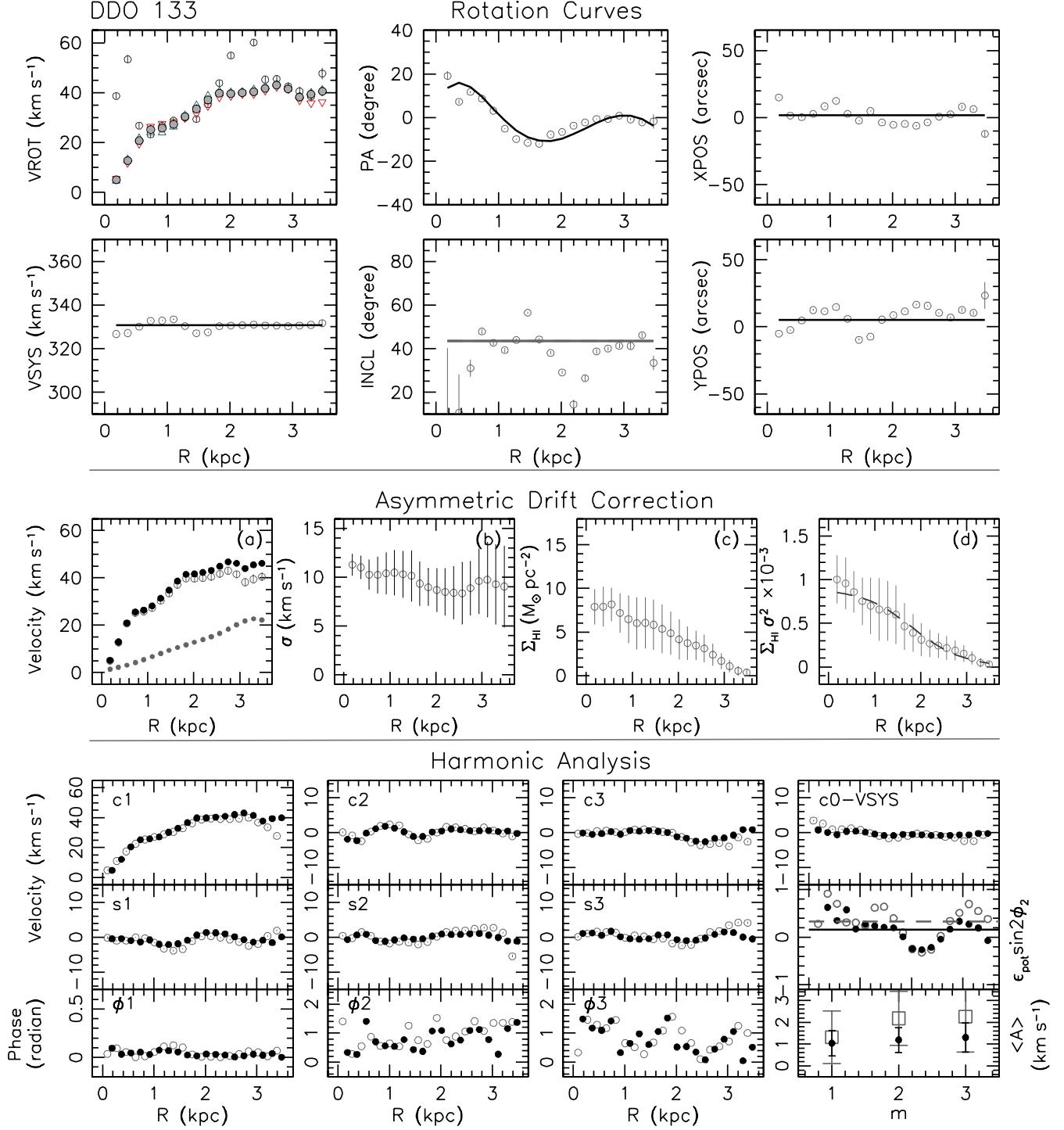
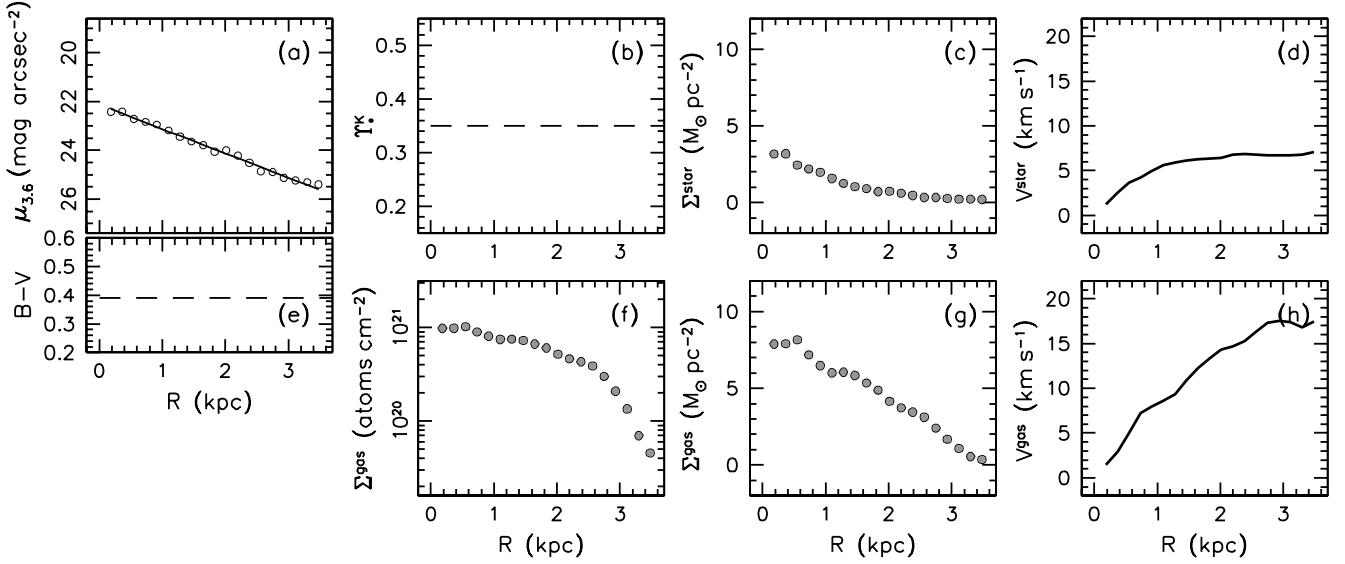


FIG. A.35.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 133. See Appendix section B for details.

DDO 133

Mass models of baryons



Disk–halo decomposition

Mass density profile

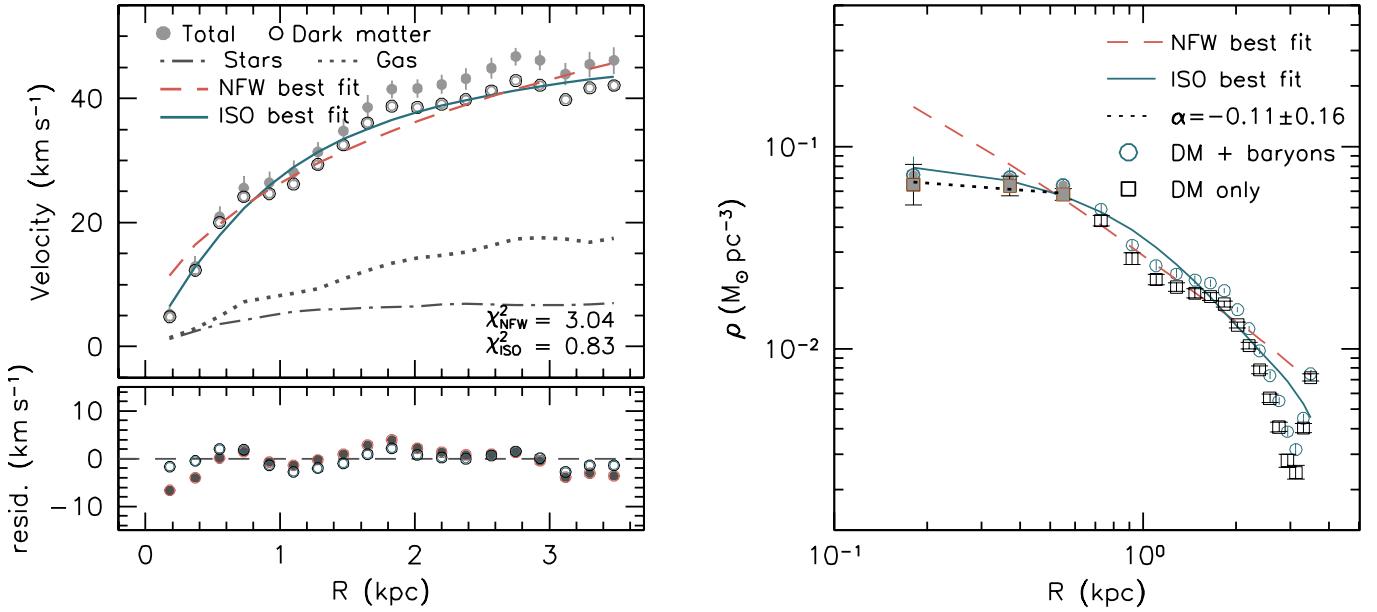


FIG. A.36.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 133. Please refer to the text in Sections 3 and 4 for full information.

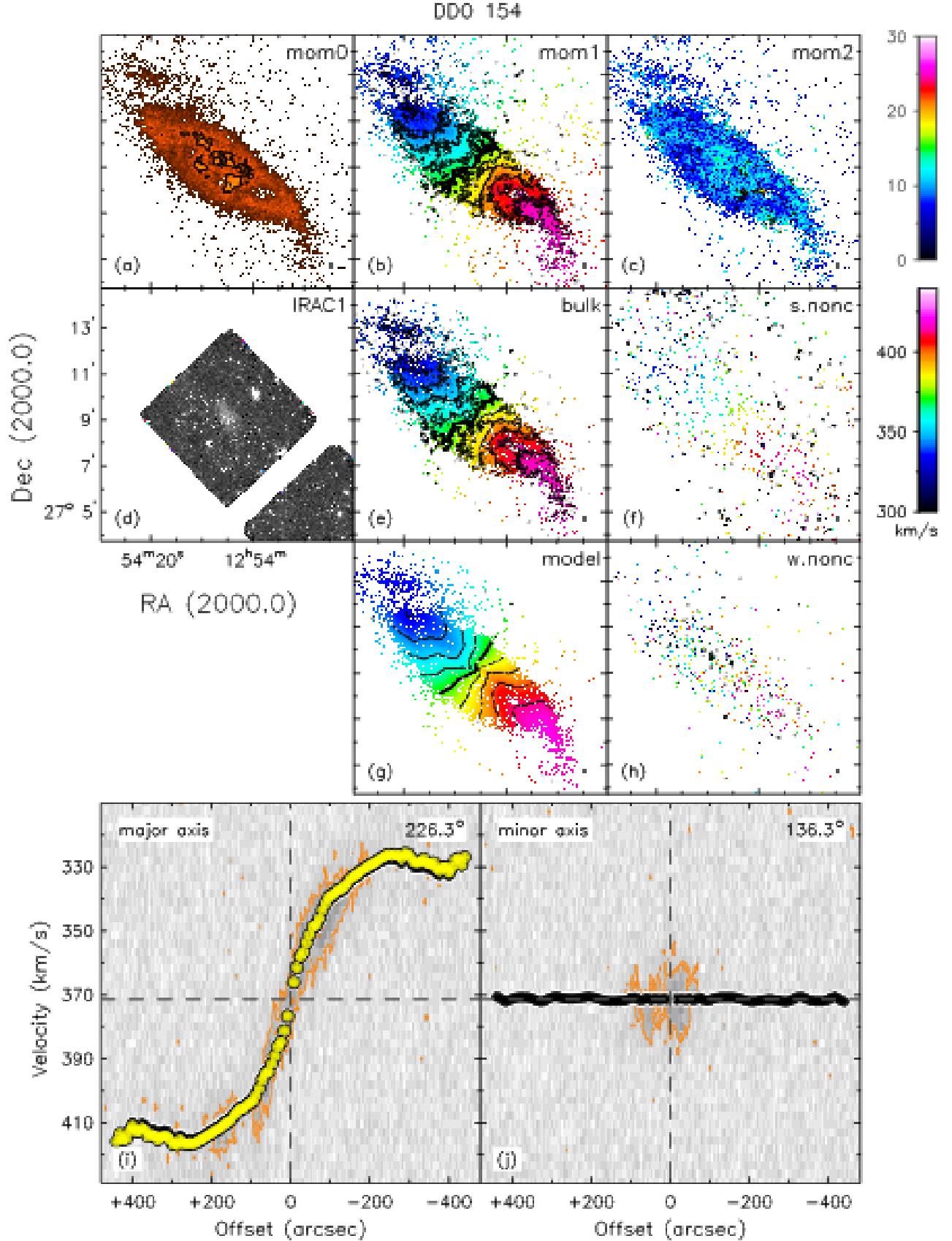


FIG. A.37.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 154. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 15 km s^{-1} . See Appendix section A for details.

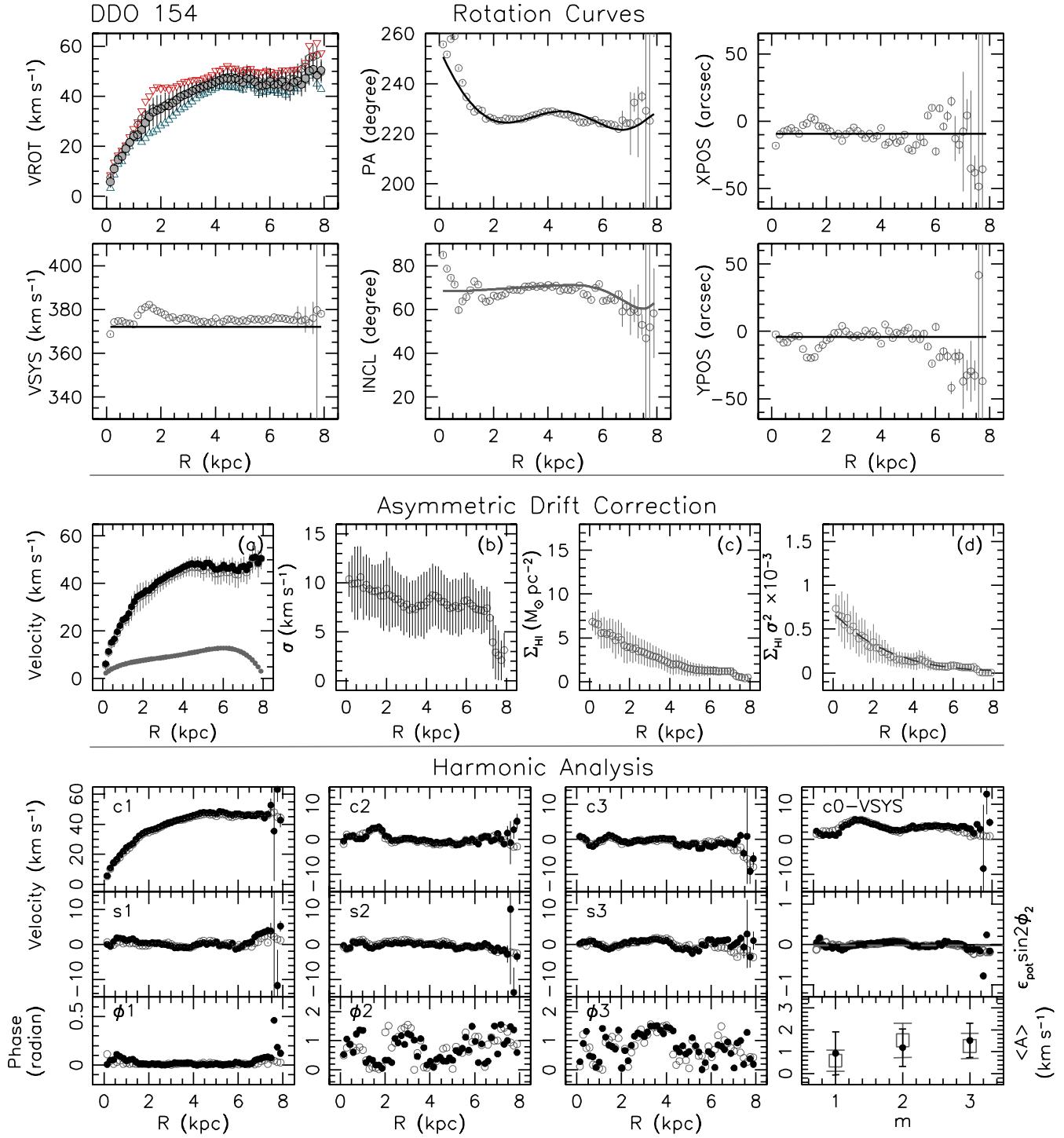


FIG. A.38.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 154. See Appendix section B for details.

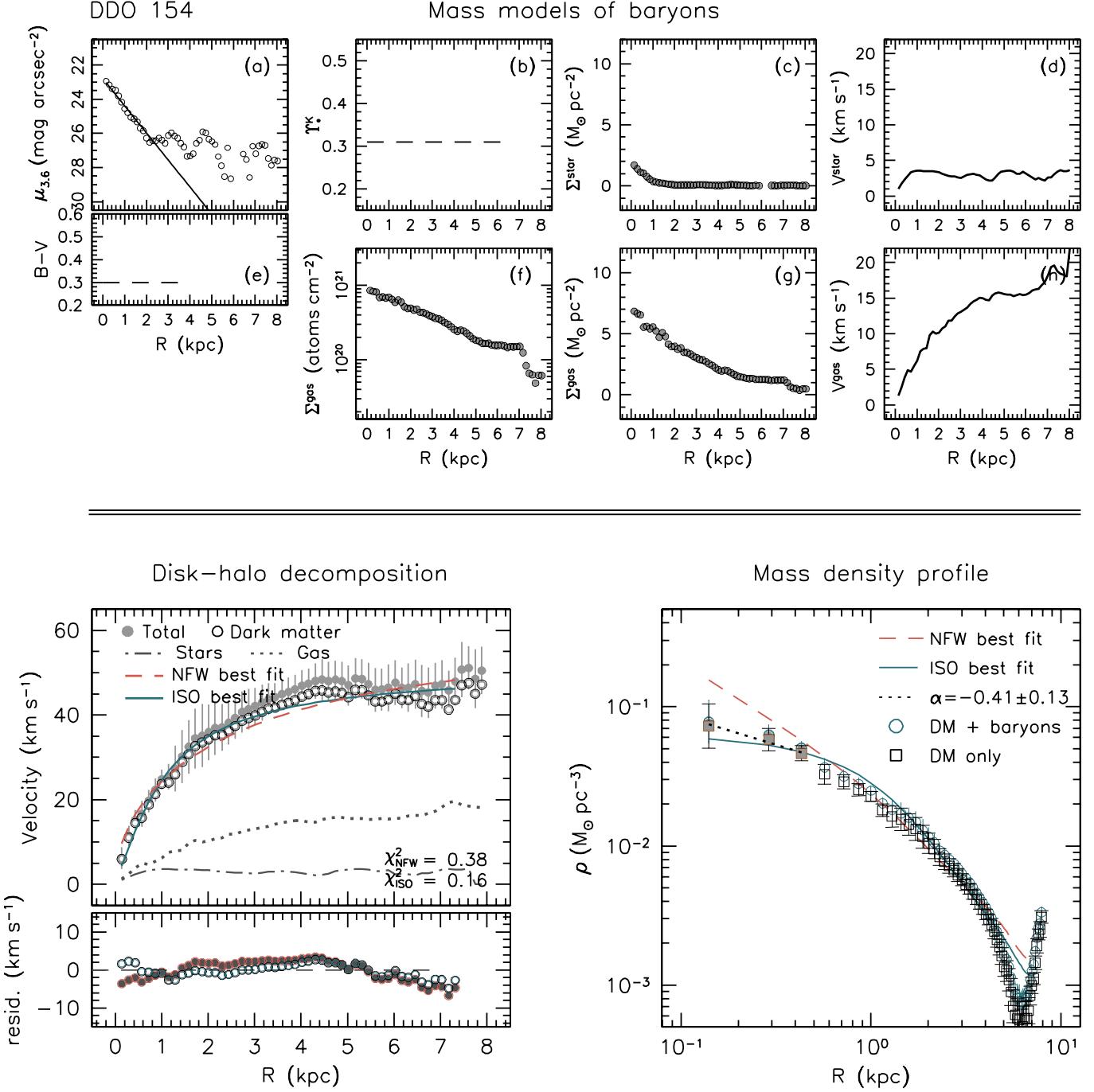


FIG. A.39.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 154. Please refer to the text in Sections 3 and 4 for full information.

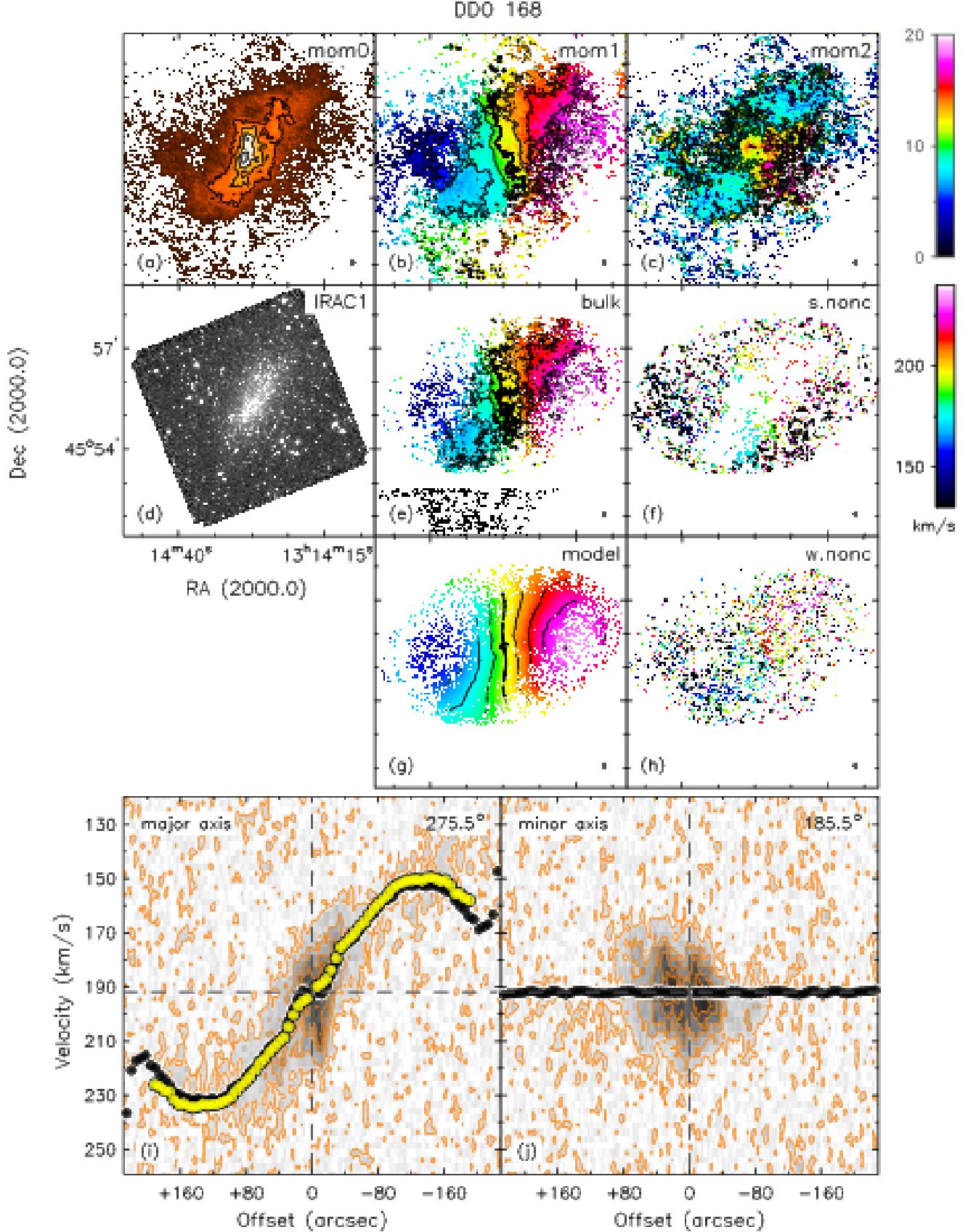


FIG. A.40.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 168. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

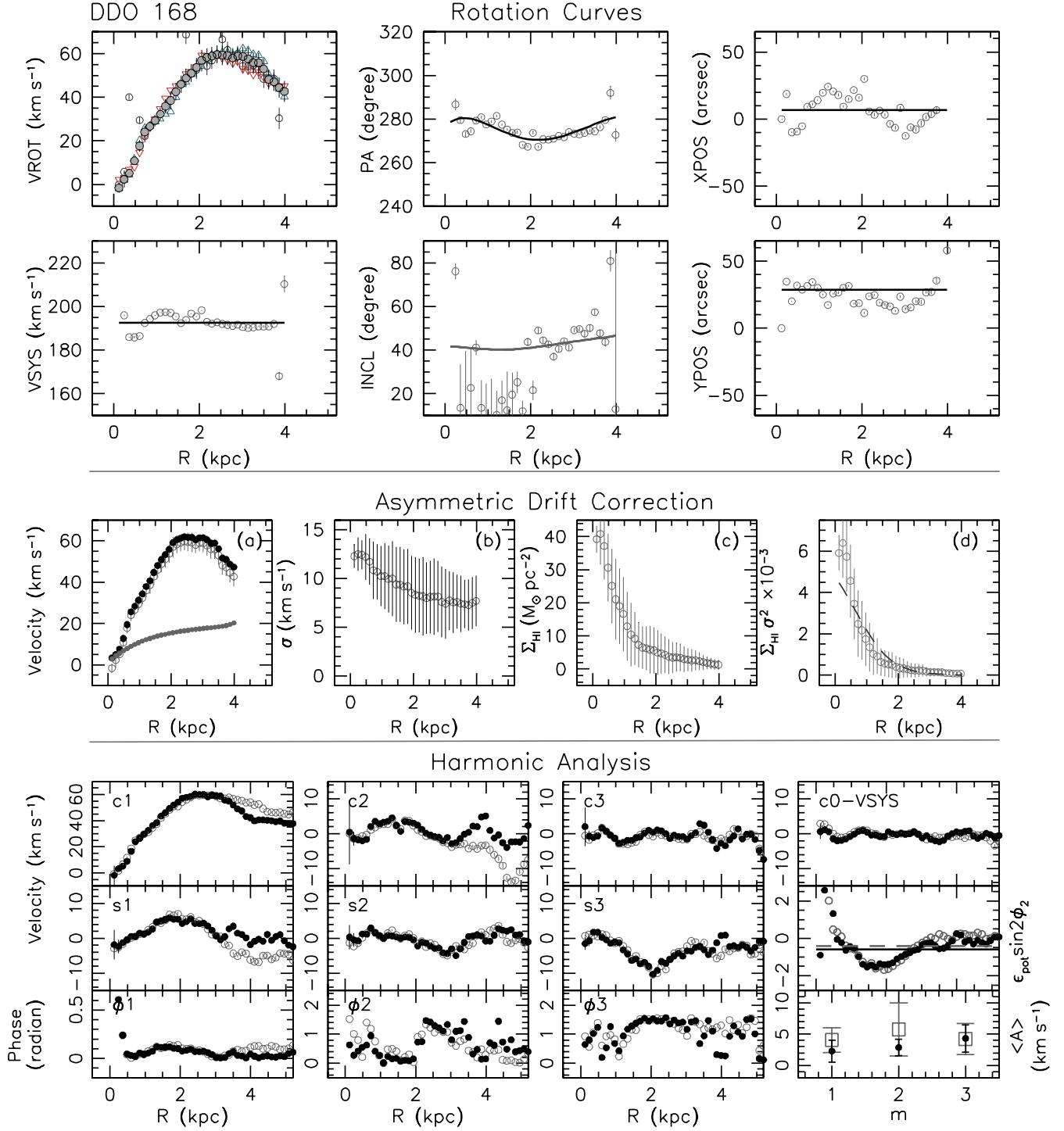
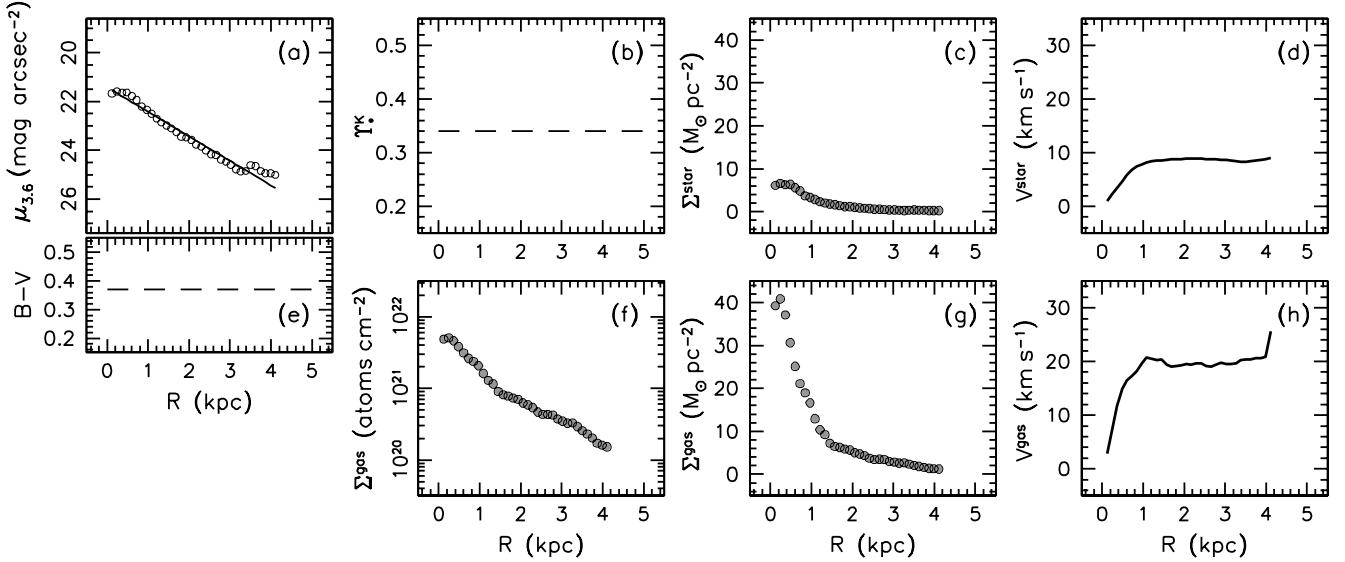


FIG. A.41.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 168. See Appendix section B for details.

DDO 168

Mass models of baryons



Disk-halo decomposition

Mass density profile

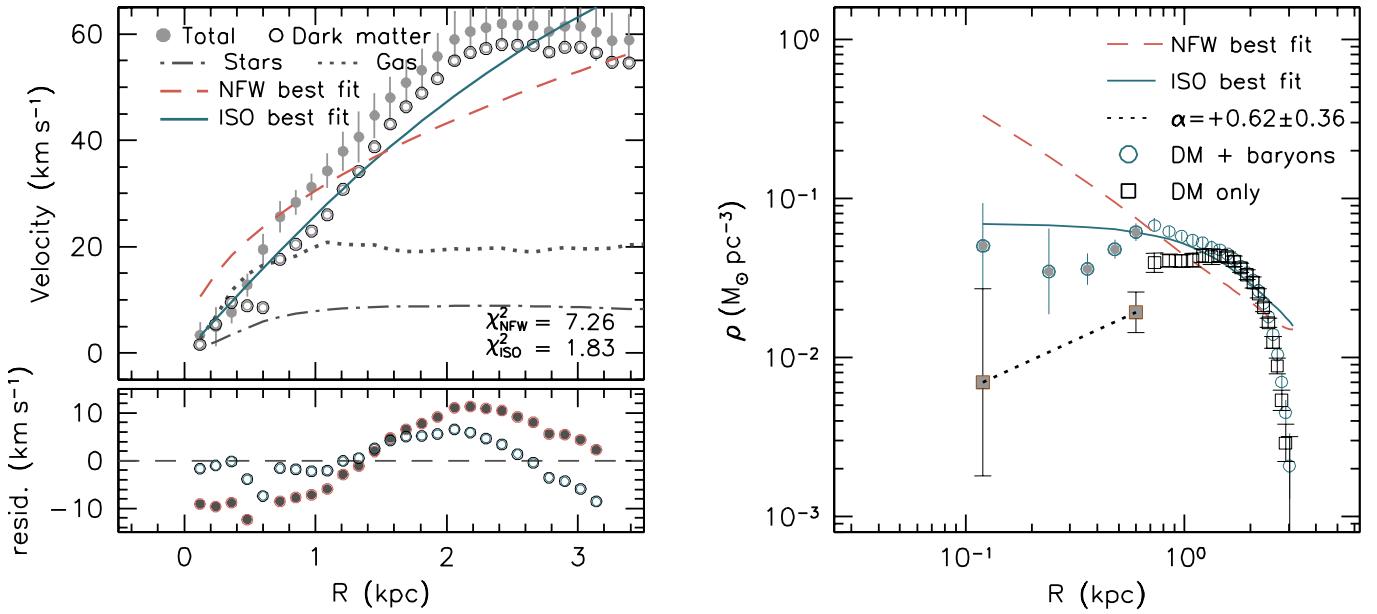


FIG. A.42.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 168. Please refer to the text in Sections 3 and 4 for full information.

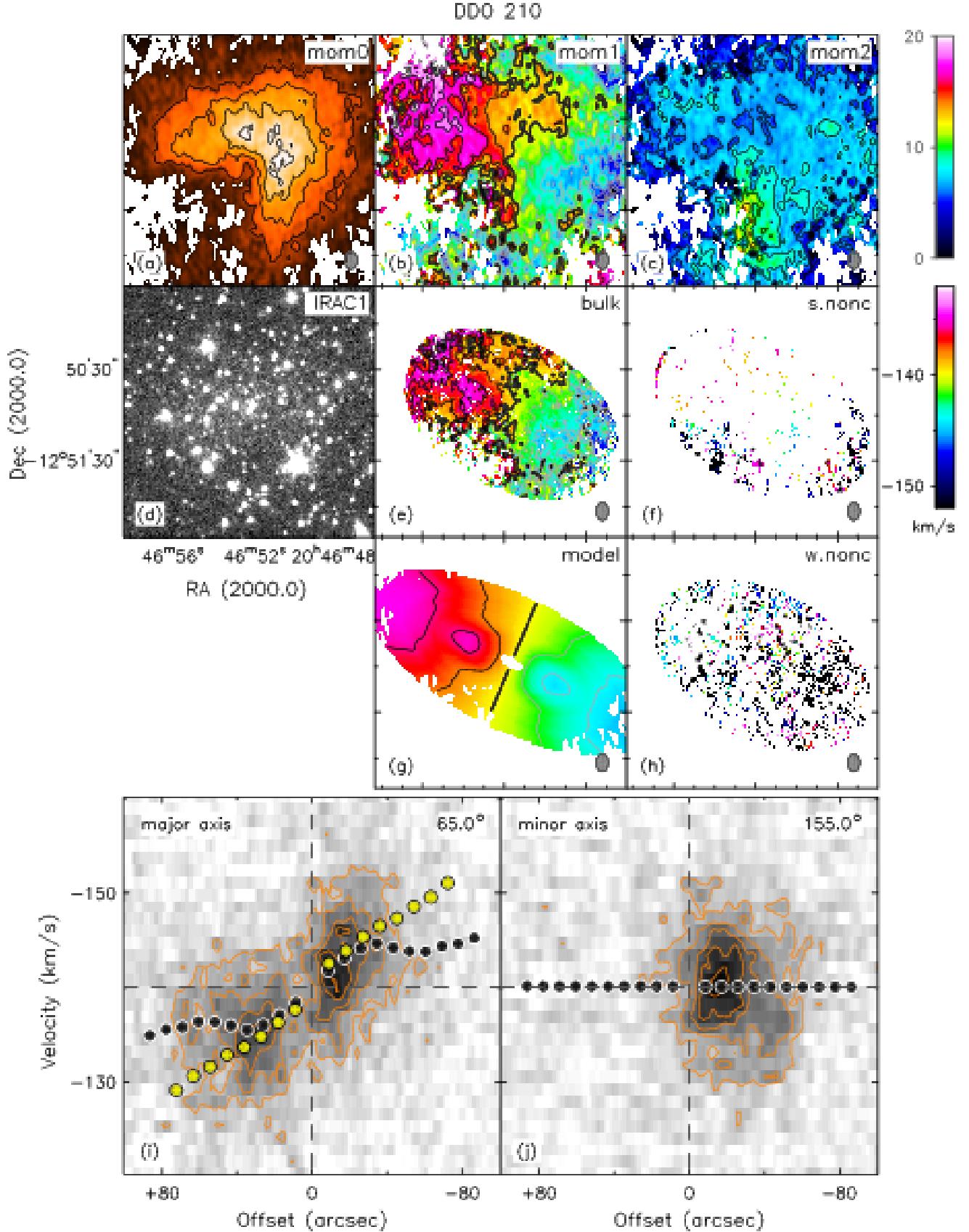


FIG. A.43.— HI data and *Spitzer* IRAC 3.6 μm image of DDO 210. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 2 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 2 km s^{-1} . See Appendix section A for details.

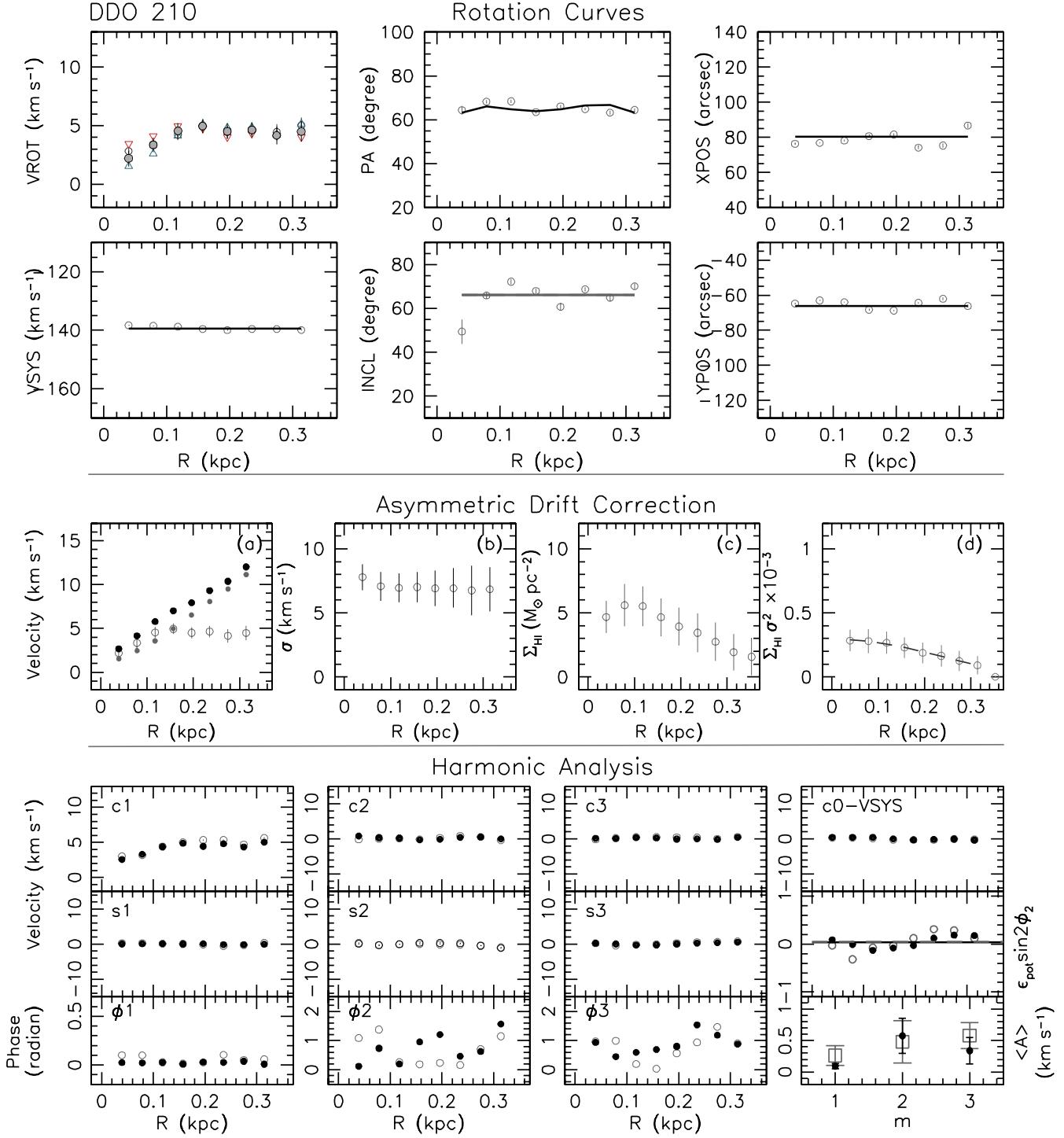


FIG. A.44.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 210. See Appendix section B for details.

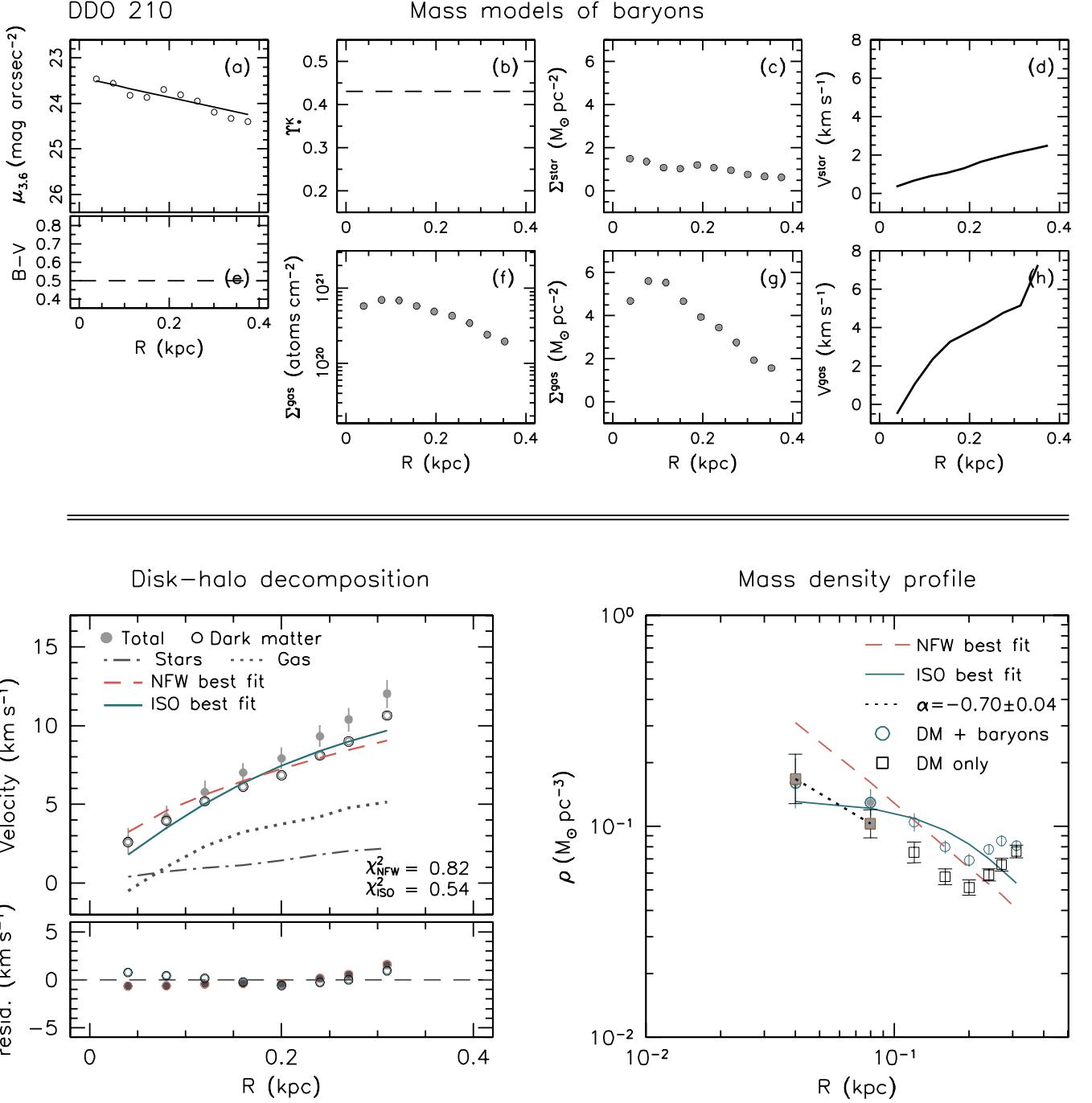


FIG. A.45.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 210. Please refer to the text in Sections 3 and 4 for full information.

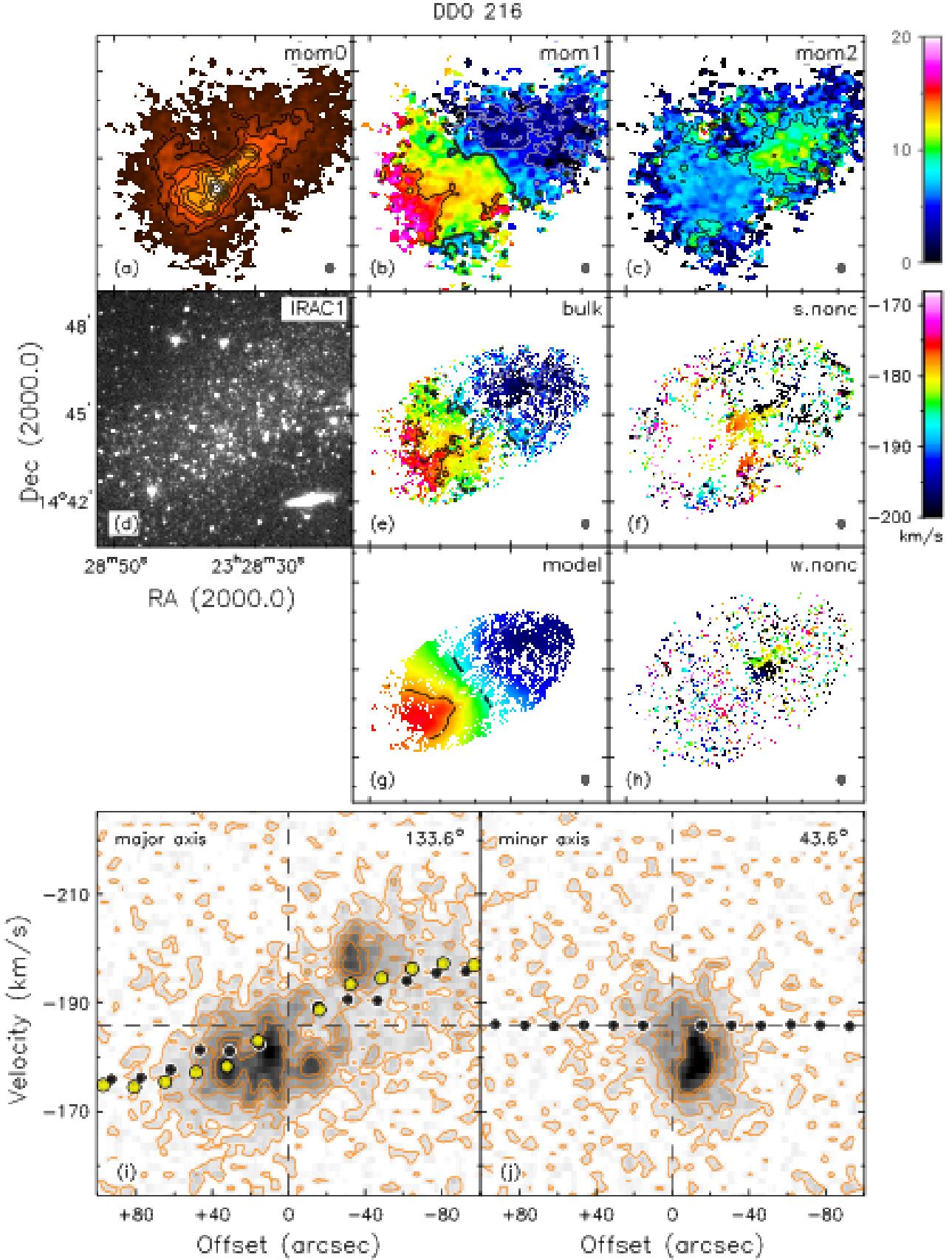


FIG. A.46.— HI data and *Spitzer* IRAC 3.6 μ m image of DDO 216. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 4 km s $^{-1}$. See Appendix section A for details.

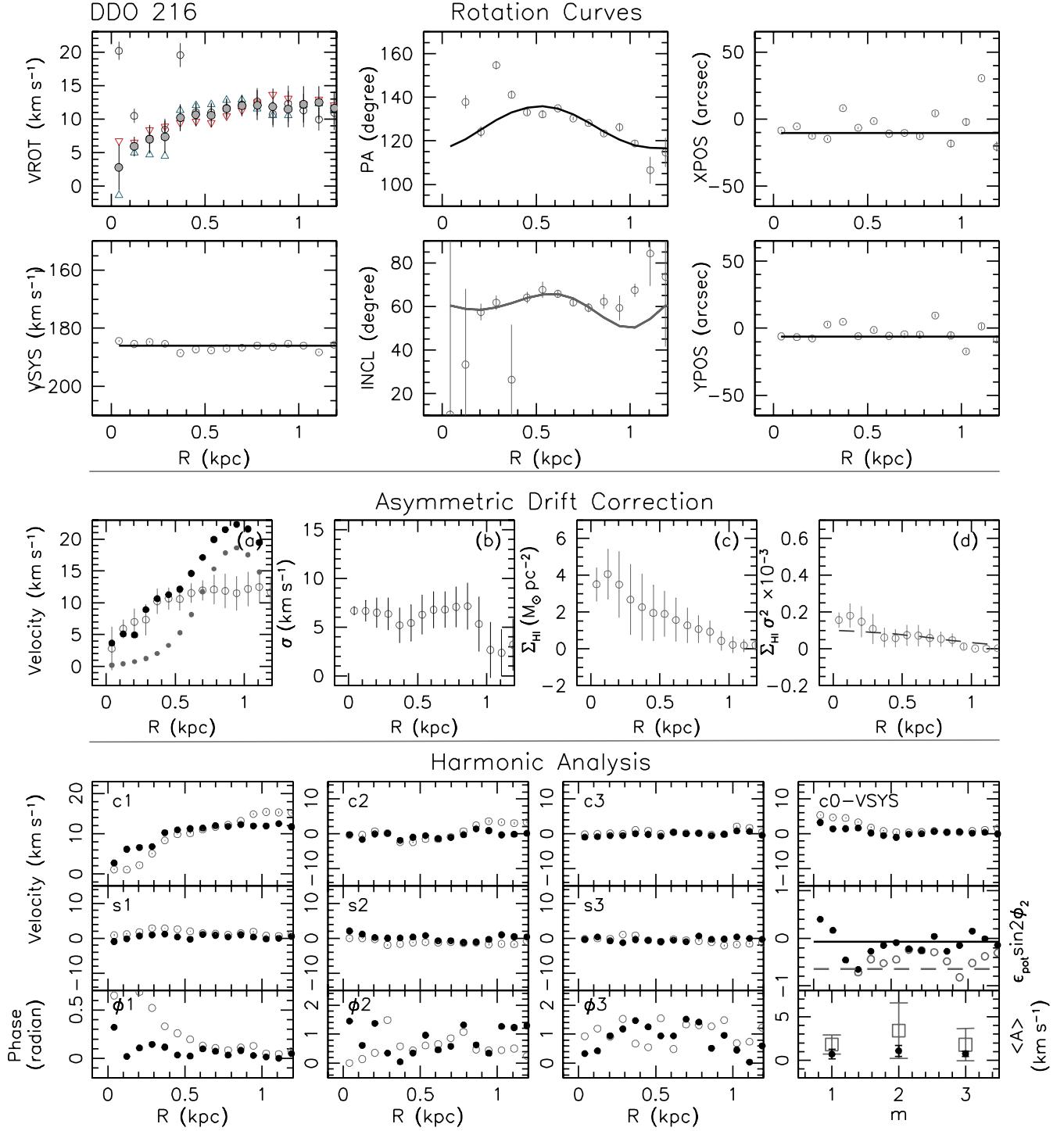
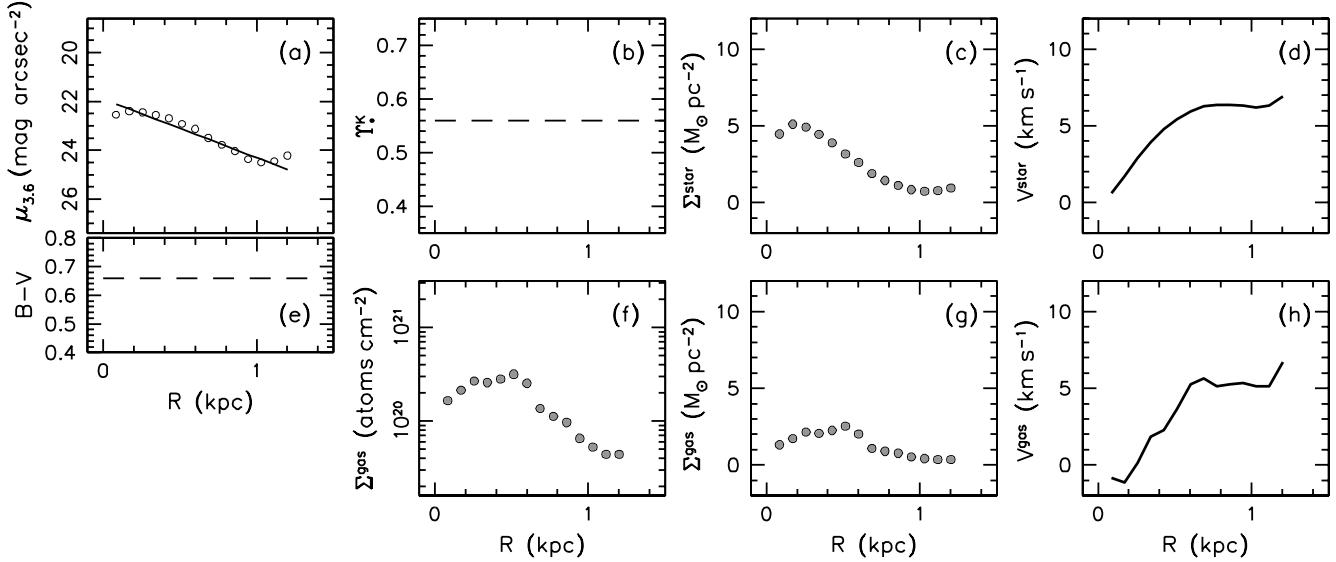


FIG. A.47.— Rotation curves, asymmetric drift correction and harmonic analysis of DDO 216. See Appendix section B for details.

DDO 216

Mass models of baryons



Disk–halo decomposition

Mass density profile

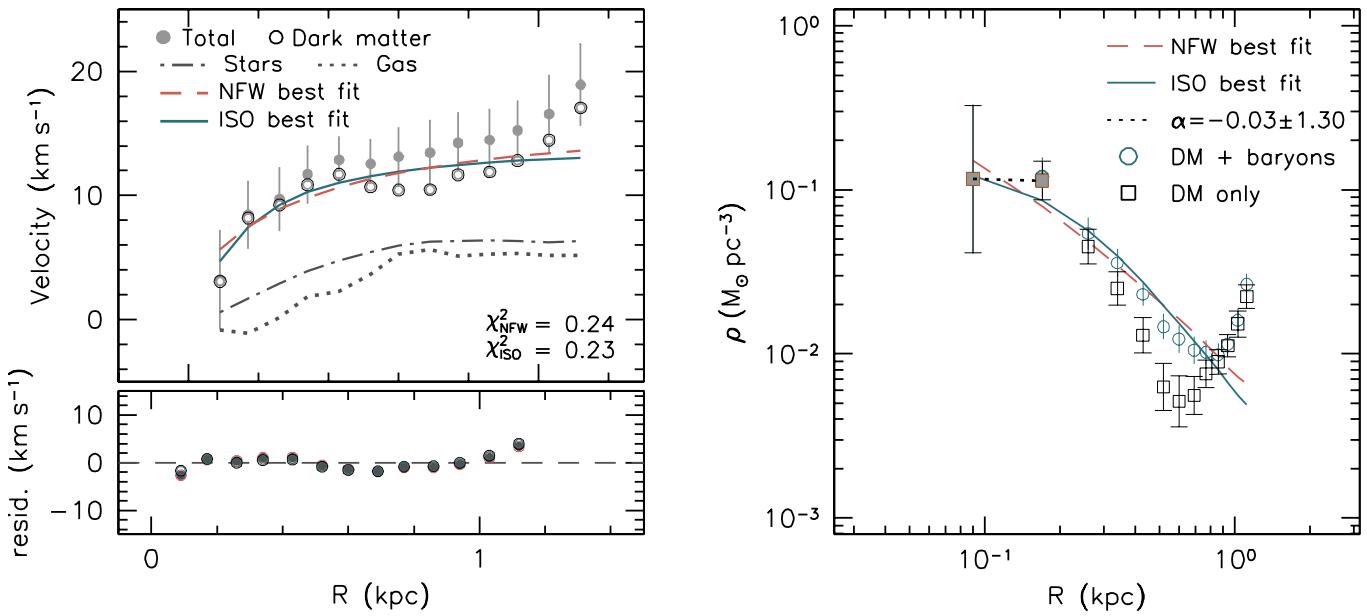


FIG. A.48.— The mass models of baryons, disk-halo decomposition and mass density profile of DDO 216. Please refer to the text in Sections 3 and 4 for full information.

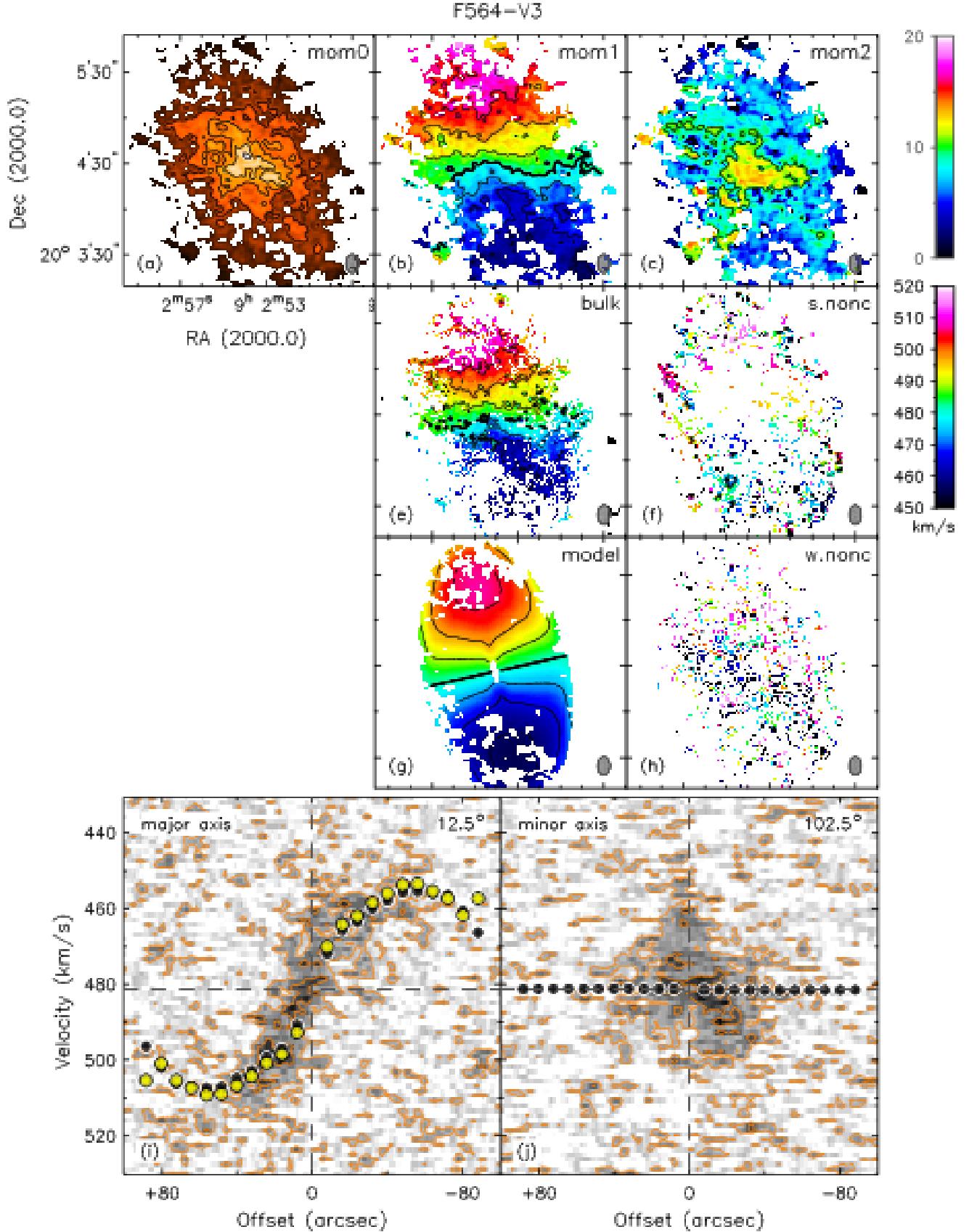


FIG. A.49.— HI data and *Spitzer* IRAC 3.6 μ m image of F564-V3. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

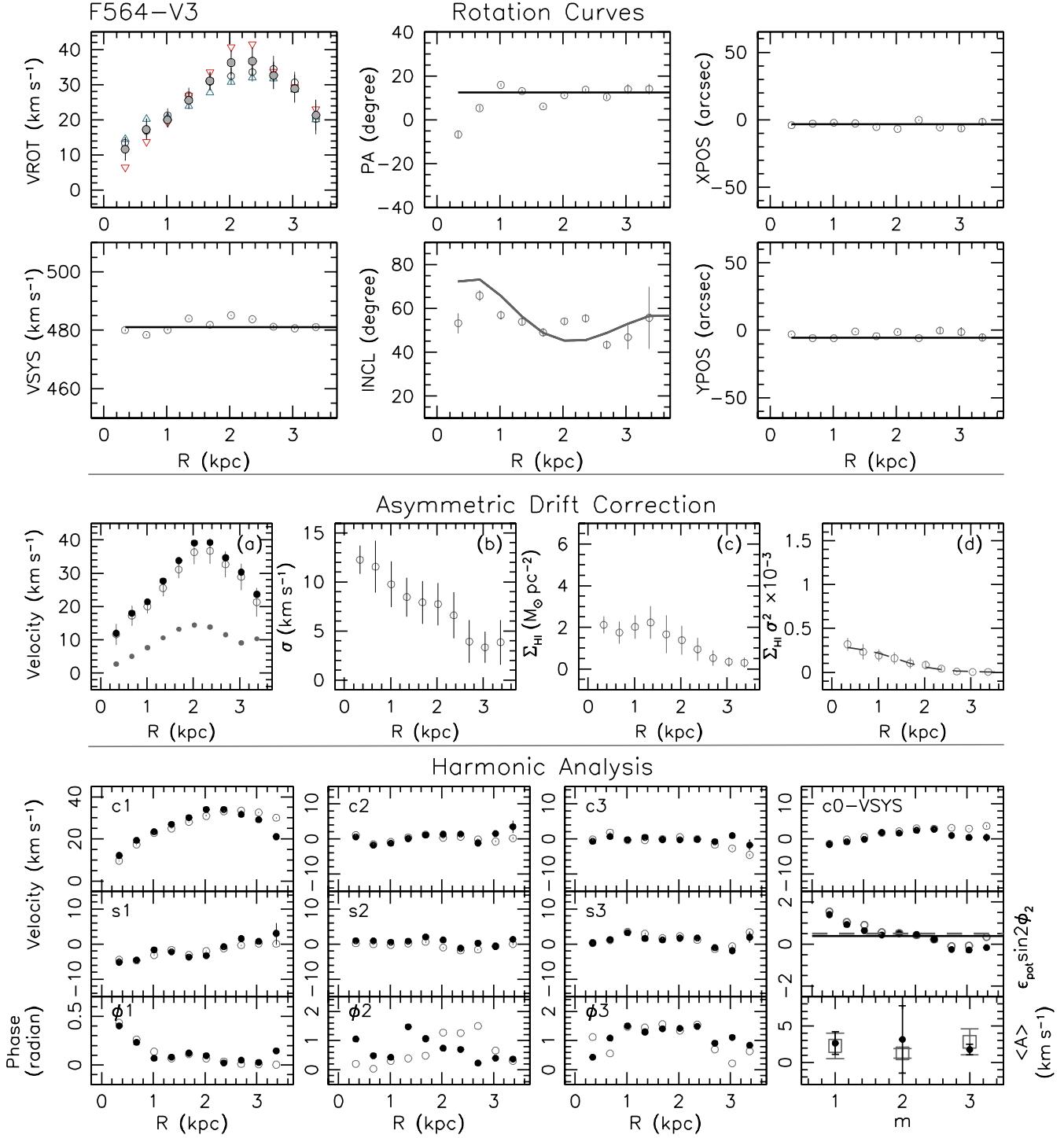


FIG. A.50.— Rotation curves, asymmetric drift correction and harmonic analysis of F564-V3. See Appendix section B for details.

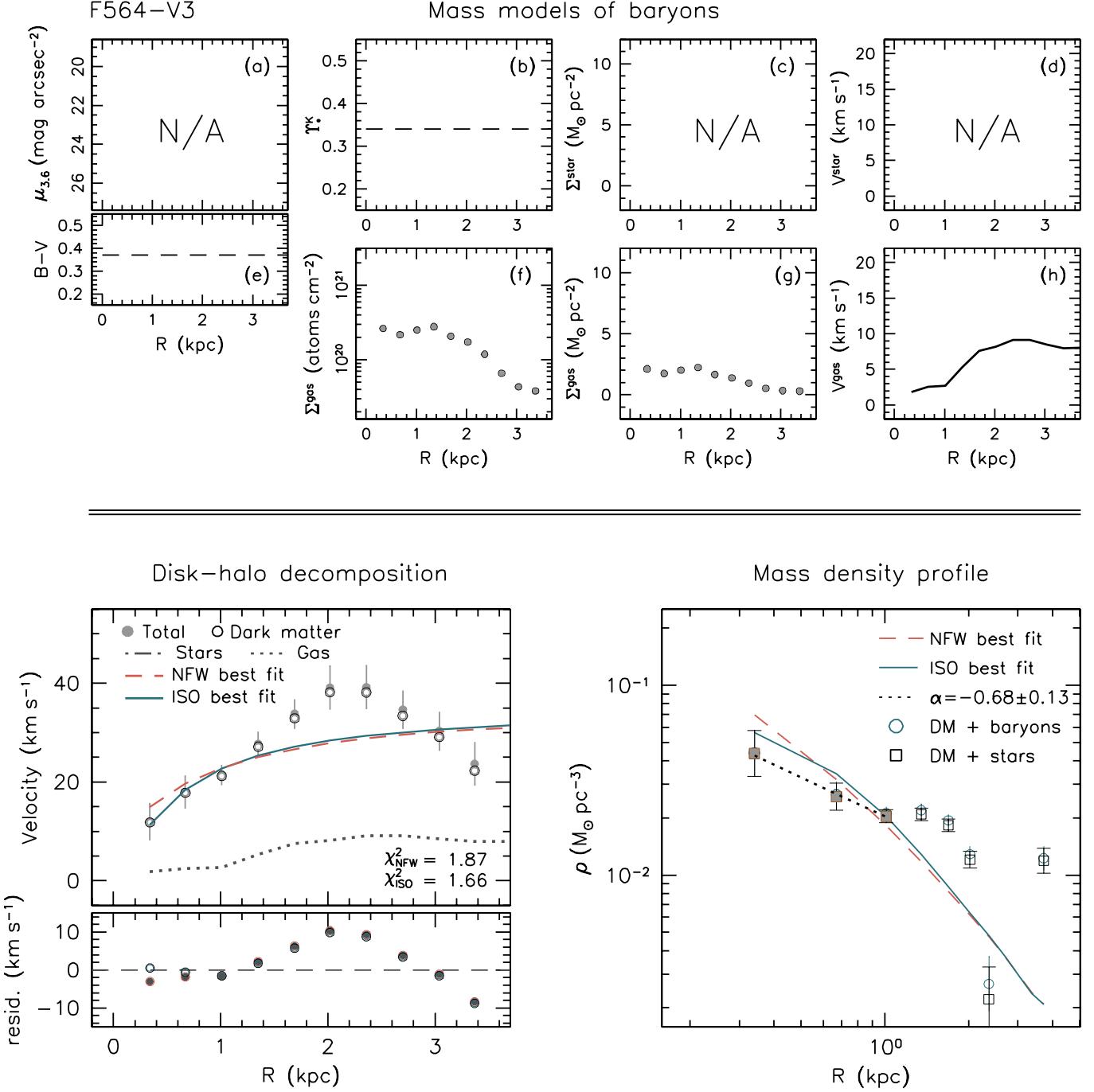


FIG. A.51.— The mass models of baryons, disk-halo decomposition and mass density profile of F564-V3. Please refer to the text in Sections 3 and 4 for full information.

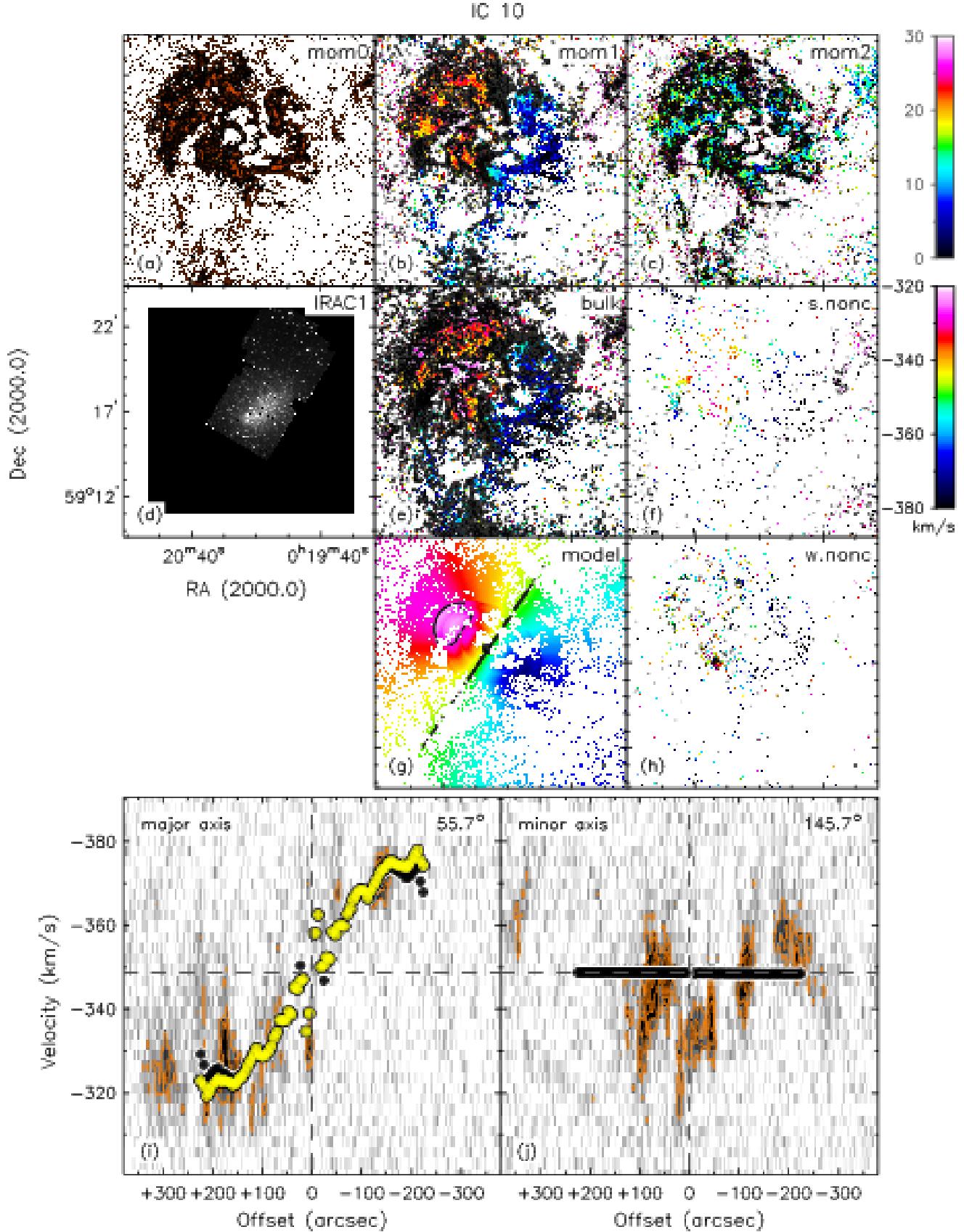


FIG. A.52.— HI data and *Spitzer* IRAC 3.6 μ m image of IC 10. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 20 km s^{-1} . Velocity dispersion contours run from 0 to 30 km s^{-1} with a spacing of 20 km s^{-1} . See Appendix section A for details.

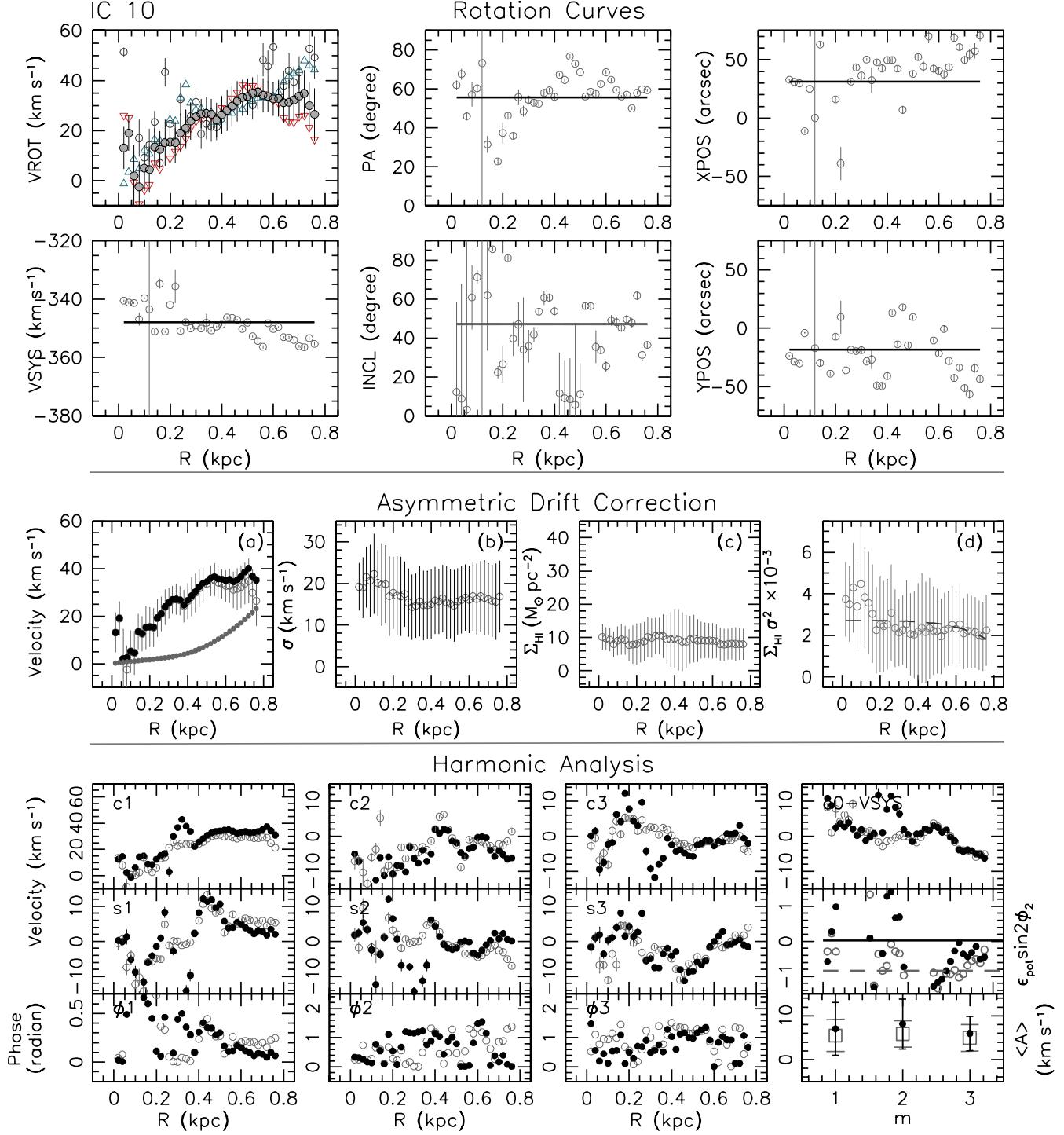
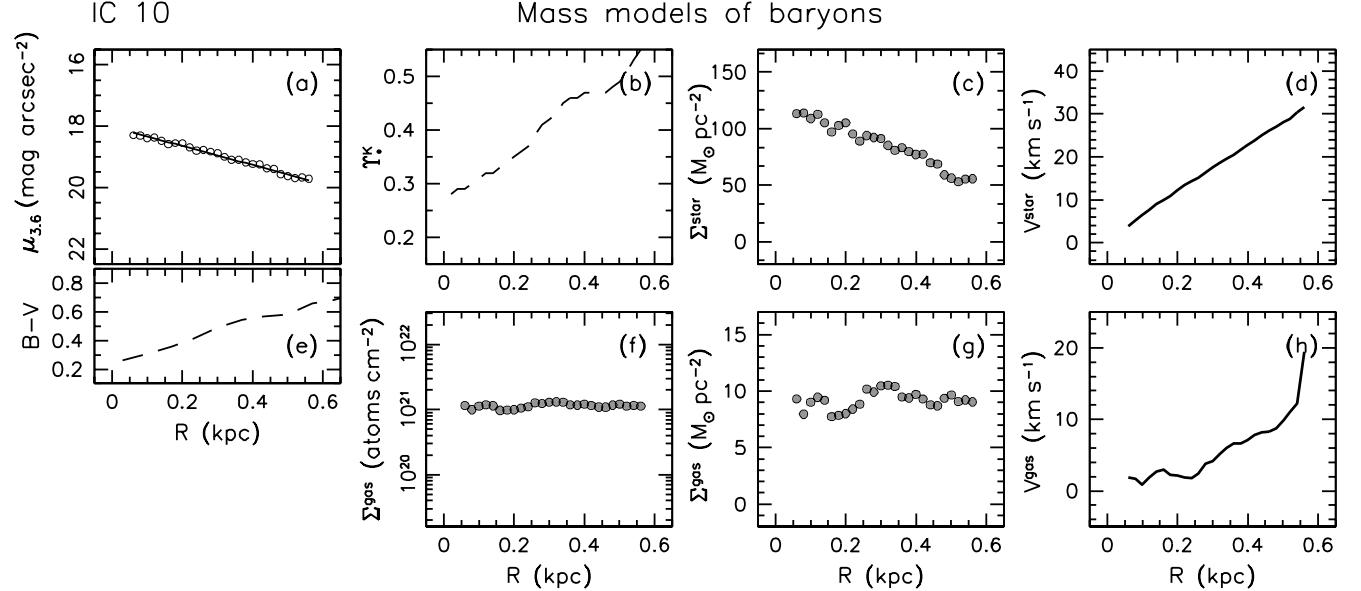
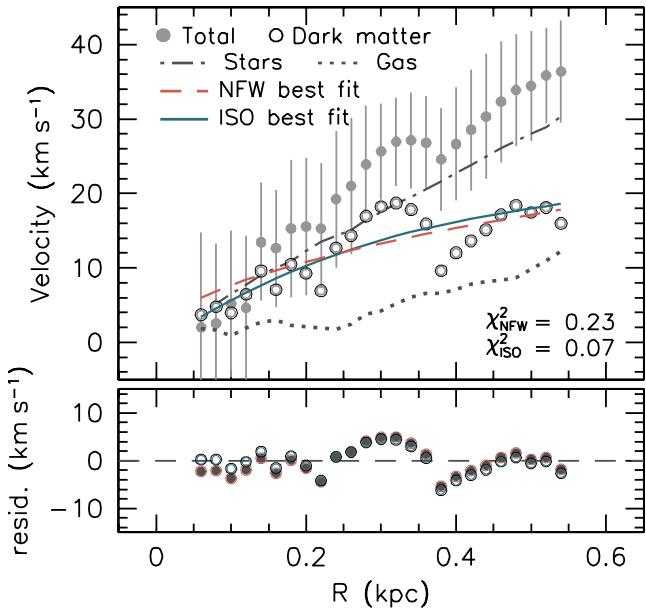


FIG. A.53.— Rotation curves, asymmetric drift correction and harmonic analysis of IC 10. See Appendix section B for details.



Disk-halo decomposition



Mass density profile

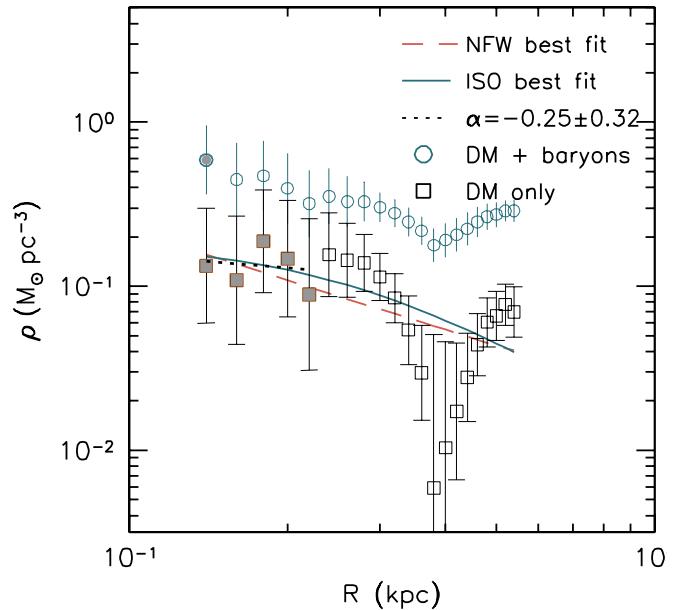


FIG. A.54.— The mass models of baryons, disk-halo decomposition and mass density profile of IC 10. Please refer to the text in Sections 3 and 4 for full information.

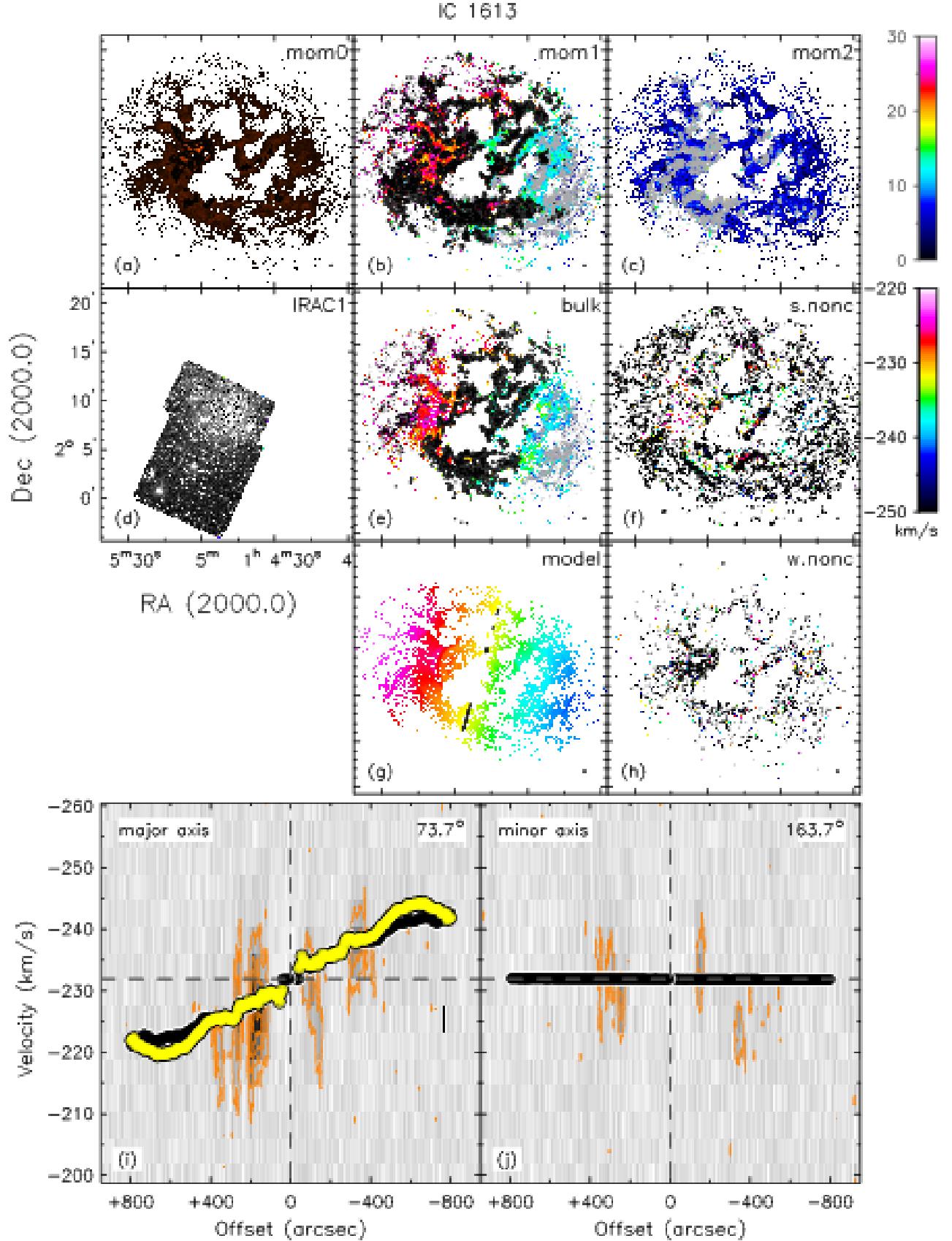


FIG. A.55.— HI data and *Spitzer* IRAC 3.6 μ m image of IC 1613. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

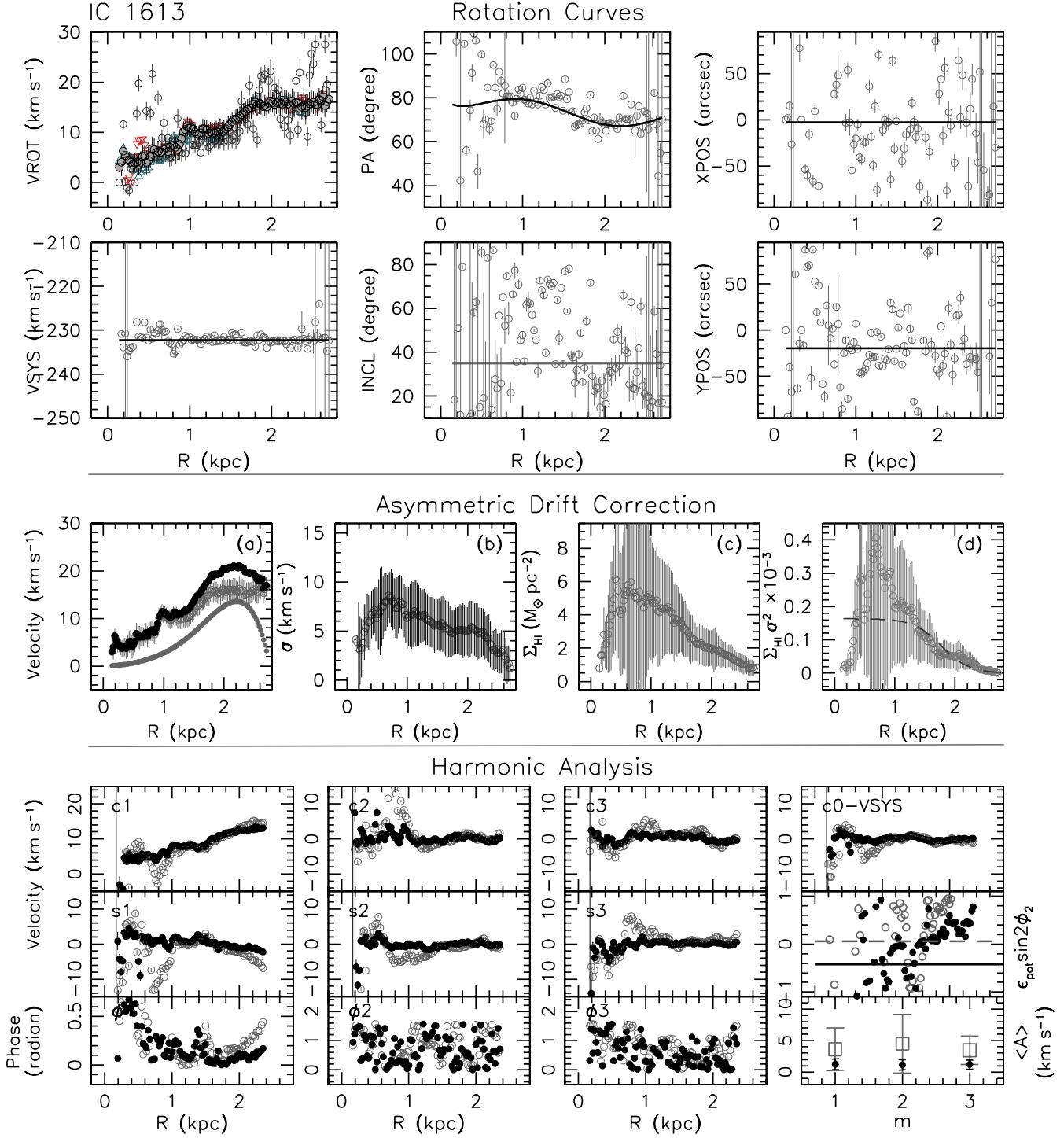


FIG. A.56.— Rotation curves, asymmetric drift correction and harmonic analysis of IC 1613. See Appendix section B for details.

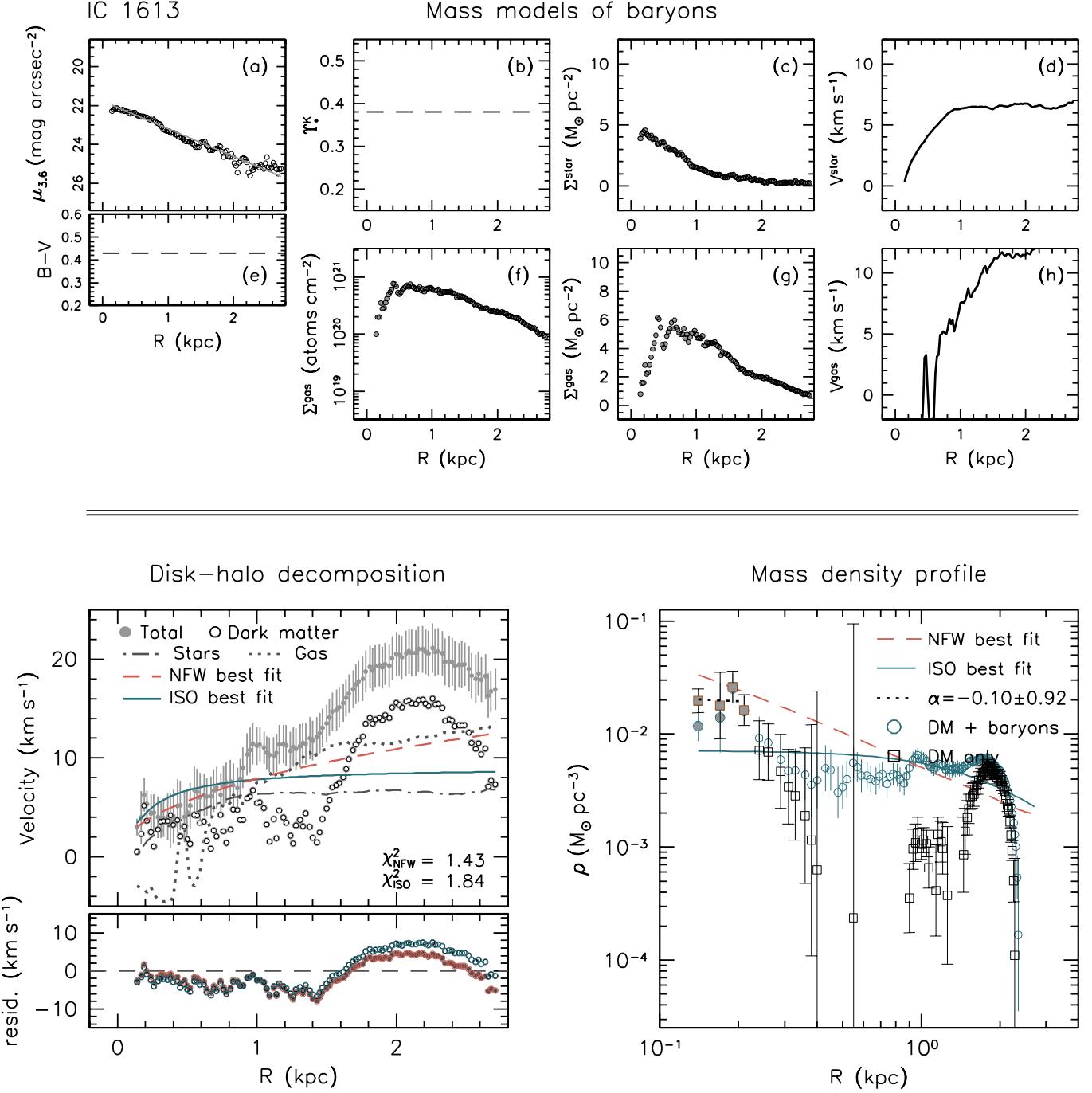


FIG. A.57.— The mass models of baryons, disk-halo decomposition and mass density profile of IC 1613. Please refer to the text in Sections 3 and 4 for full information.

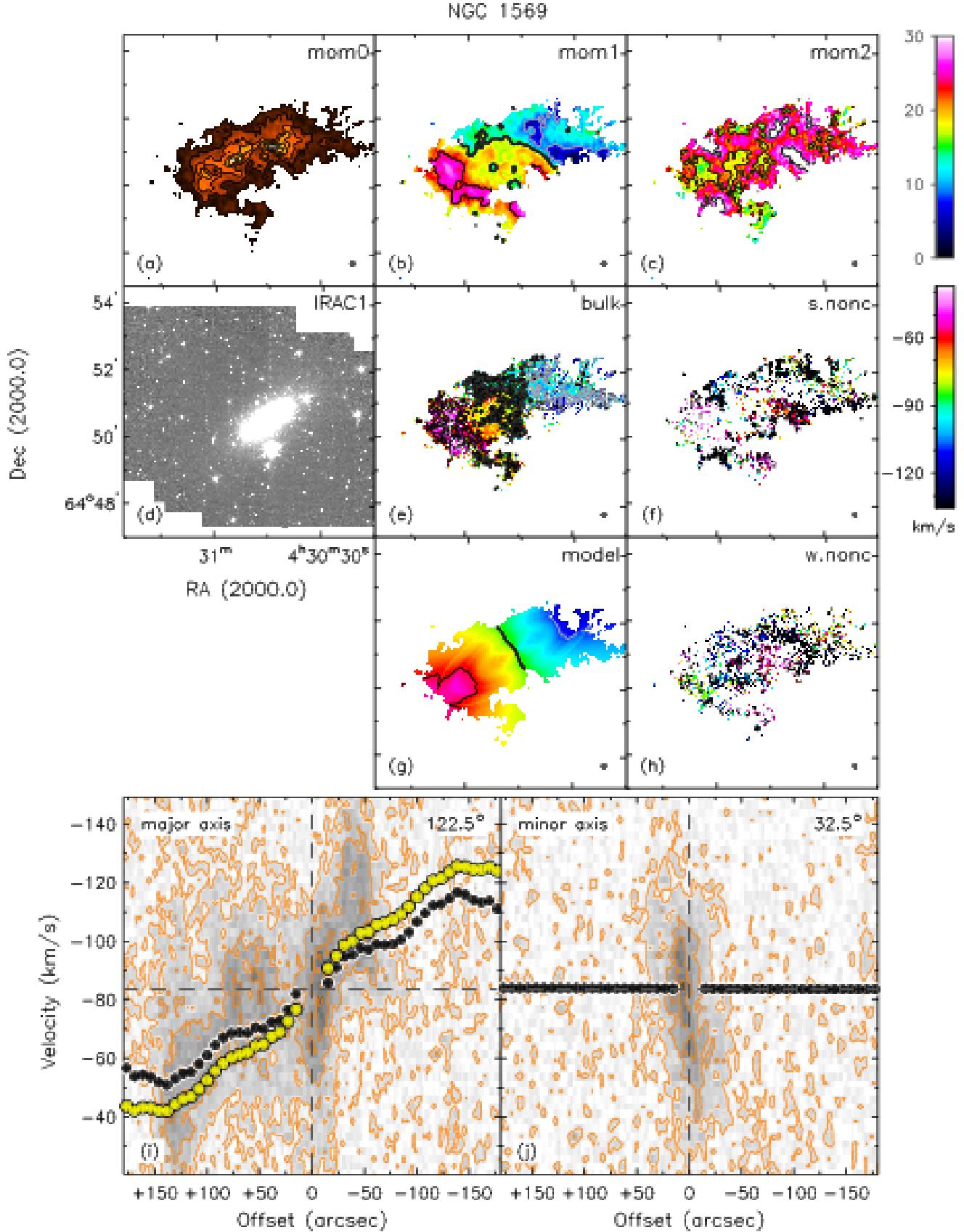


FIG. A.58.— HI data and *Spitzer* IRAC 3.6 μ m image of NGC 1569. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 25 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

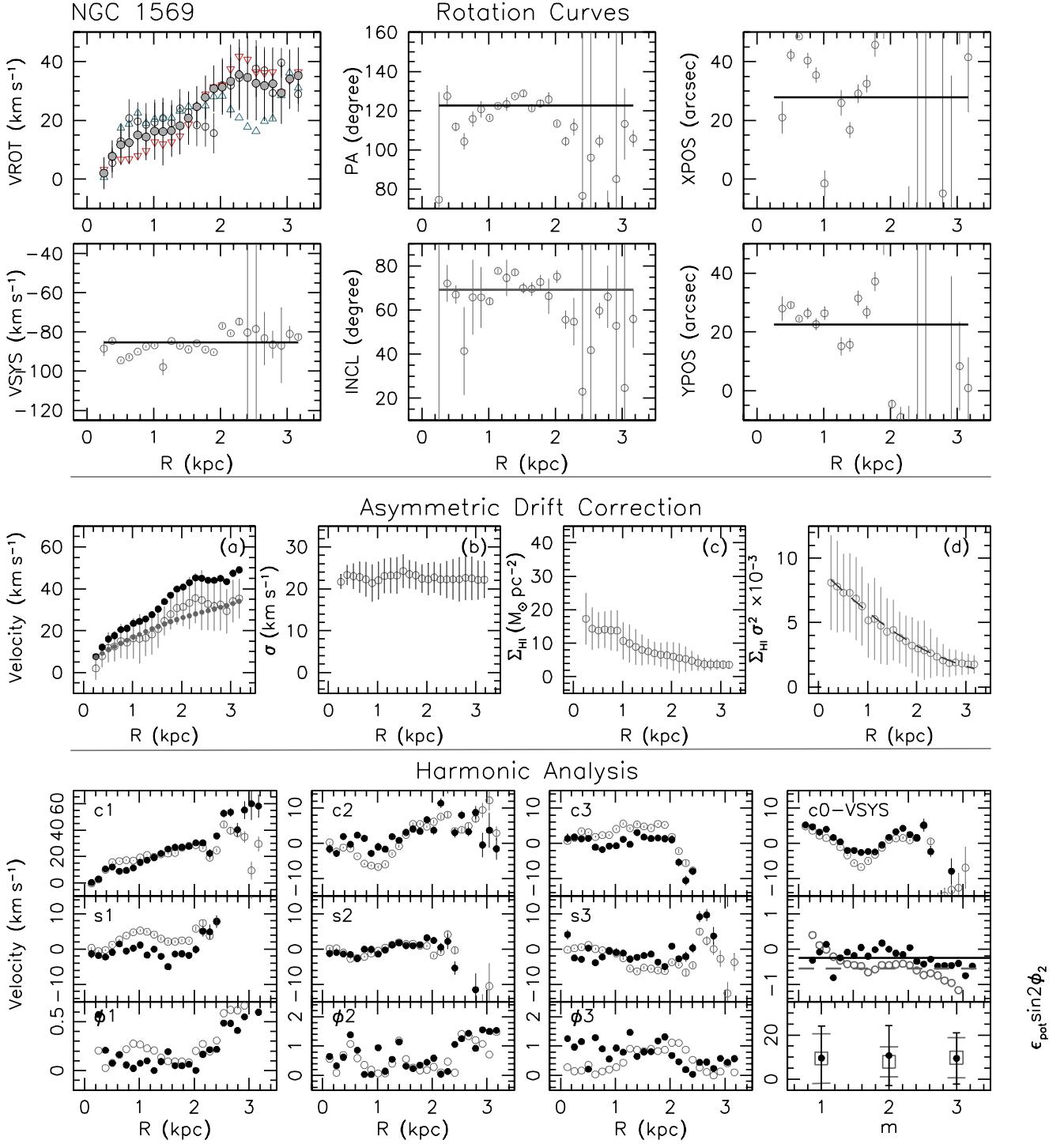
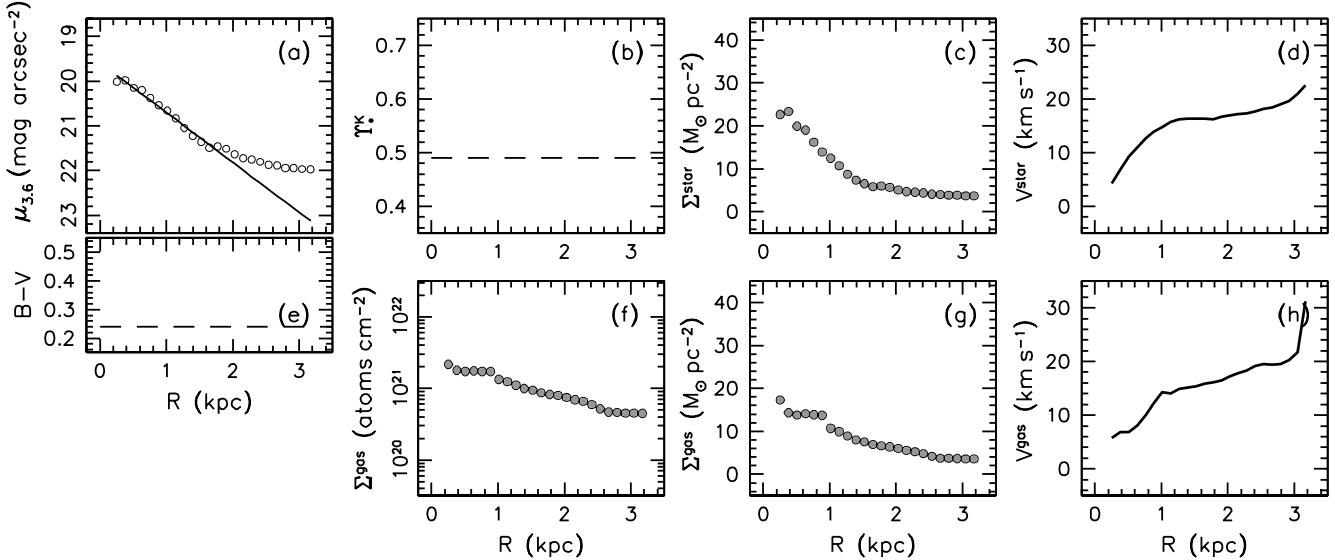


FIG. A.59.— Rotation curves, asymmetric drift correction and harmonic analysis of NGC 1569. See Appendix section B for details.

NGC 1569

Mass models of baryons



Disk–halo decomposition

Mass density profile

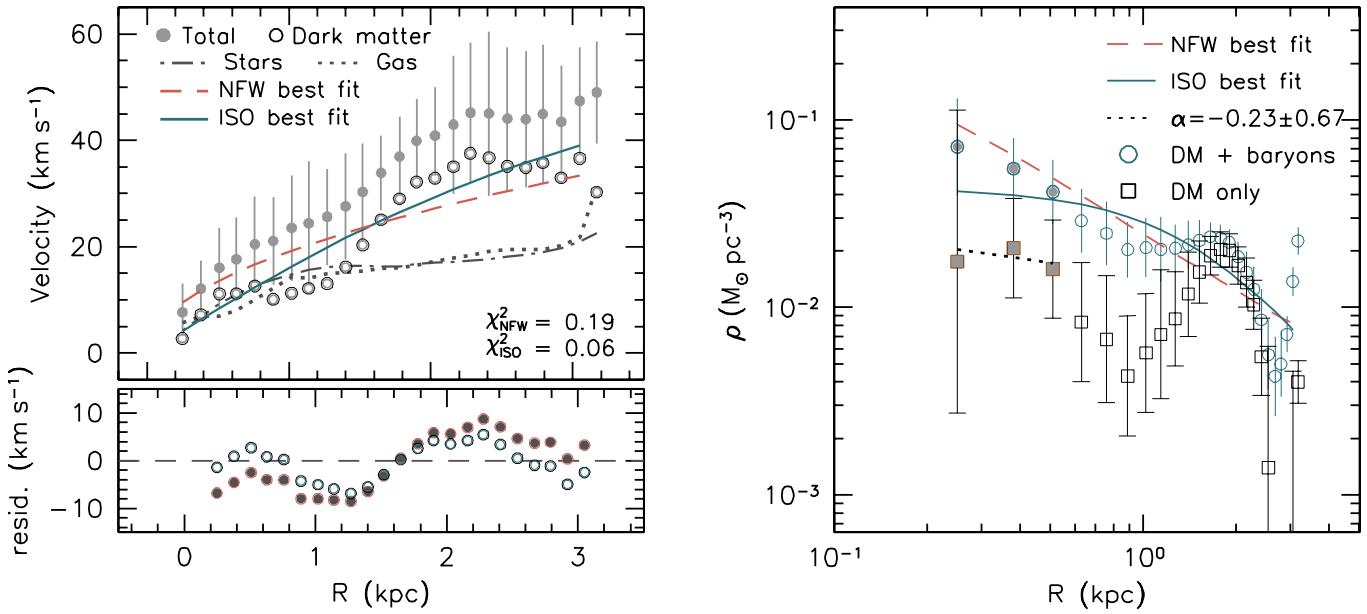


FIG. A.60.— The mass models of baryons, disk-halo decomposition and mass density profile of NGC 1569. Please refer to the text in Sections 3 and 4 for full information.

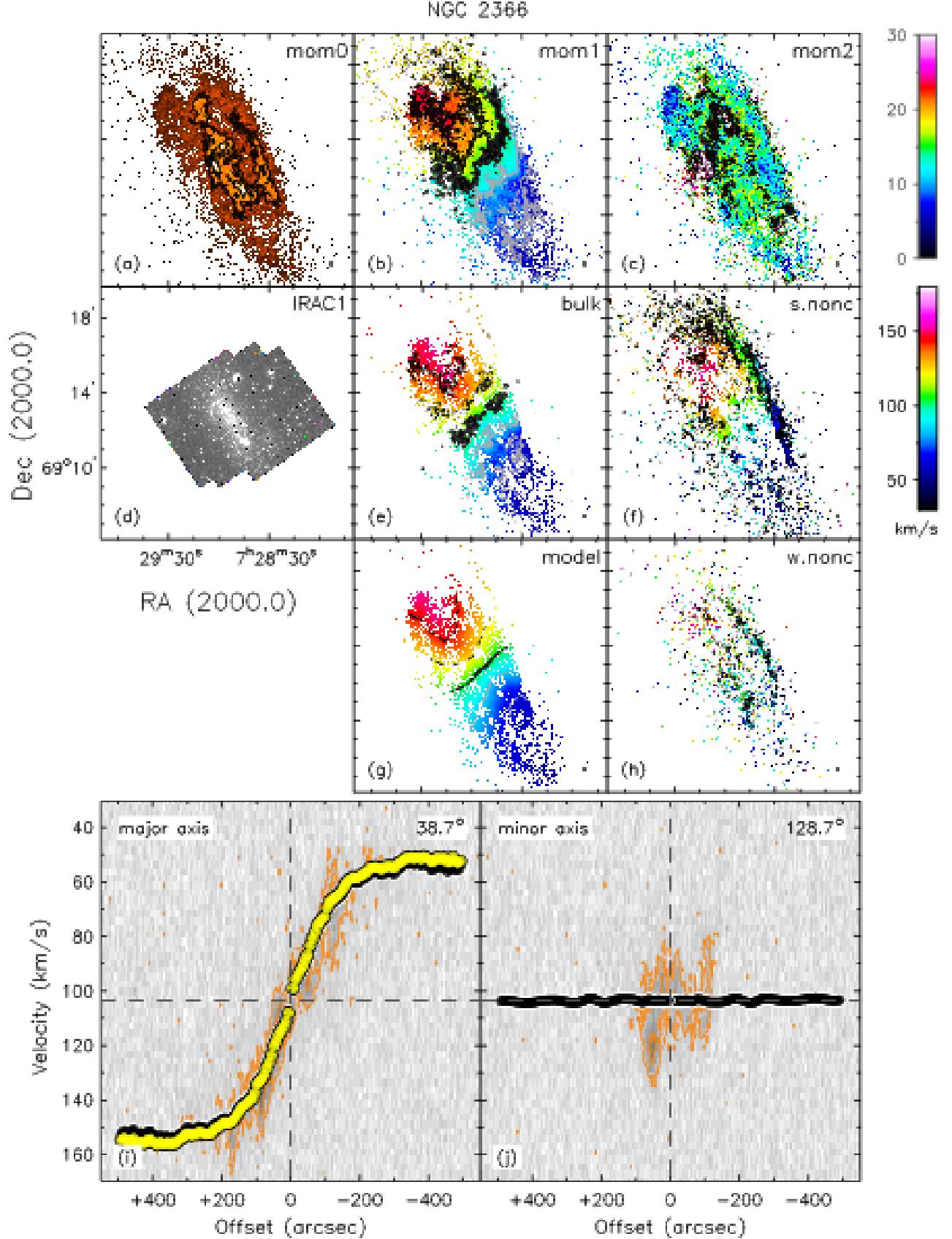


FIG. A.61.— HI data and *Spitzer* IRAC 3.6 μ m image of NGC 2366. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 20 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 20 km s^{-1} . See Appendix section A for details.

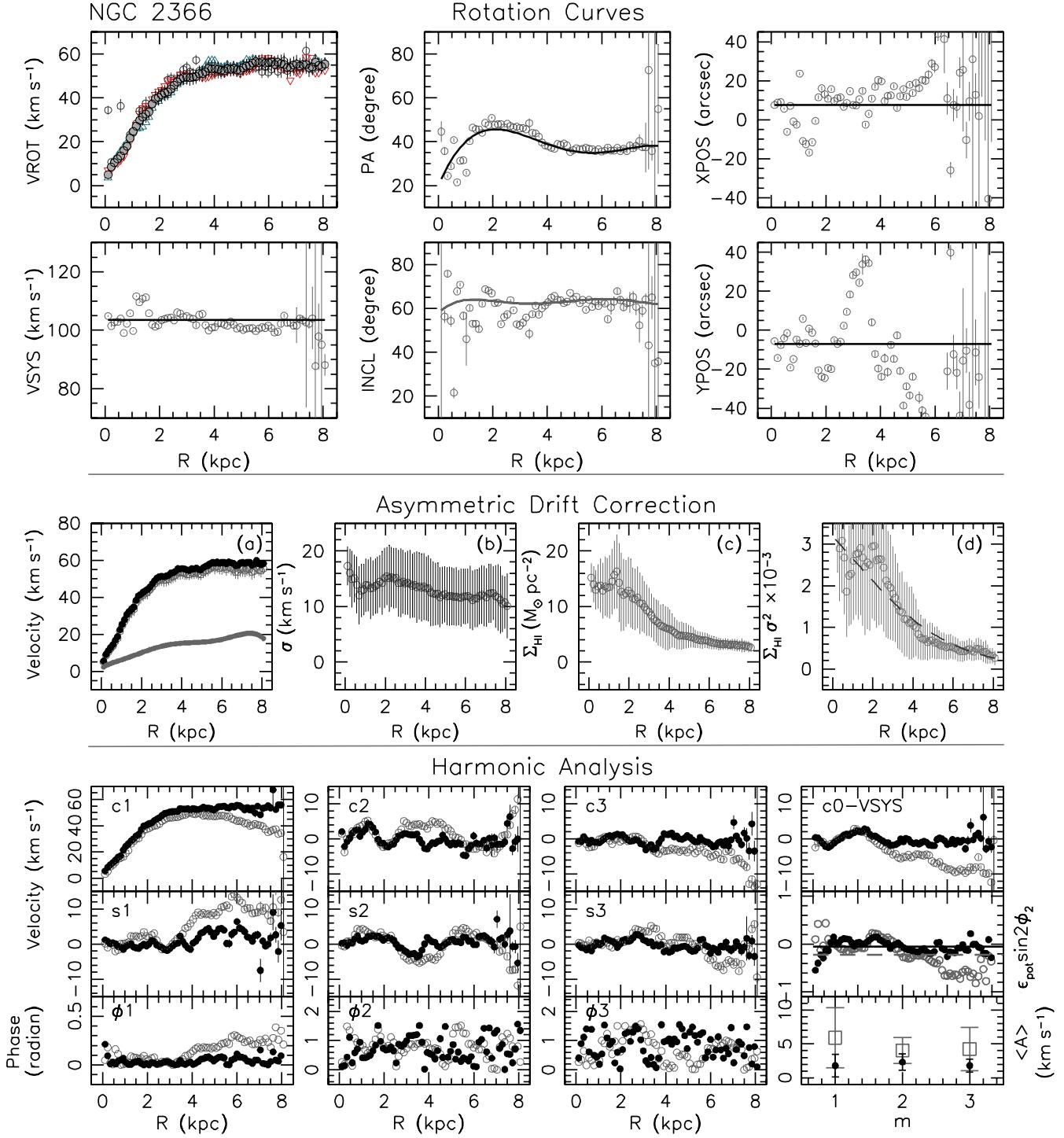


FIG. A.62.— Rotation curves, asymmetric drift correction and harmonic analysis of NGC 2366. See Appendix section B for details.

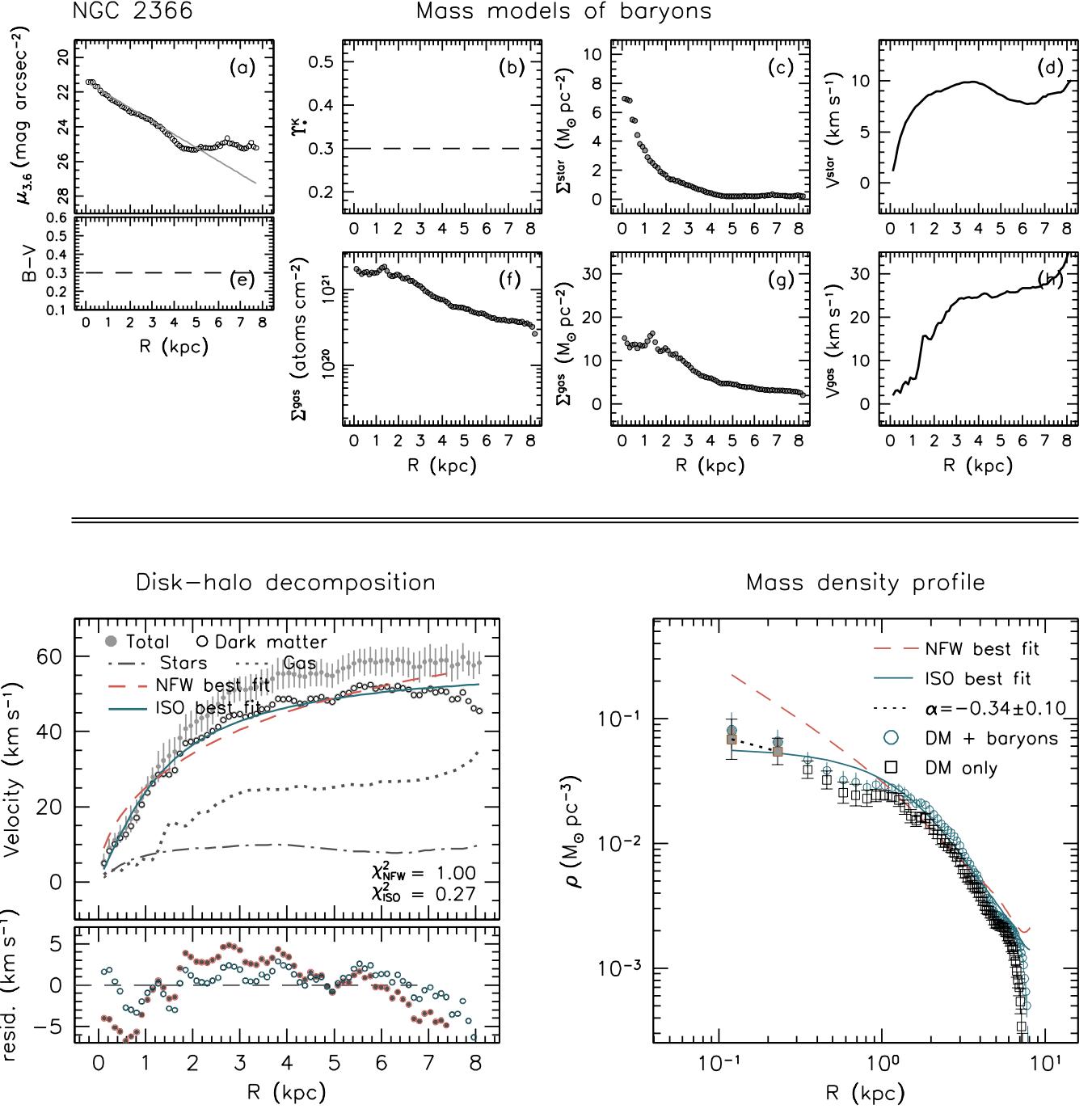


FIG. A.63.— The mass models of baryons, disk-halo decomposition and mass density profile of NGC 2366. Please refer to the text in Sections 3 and 4 for full information.

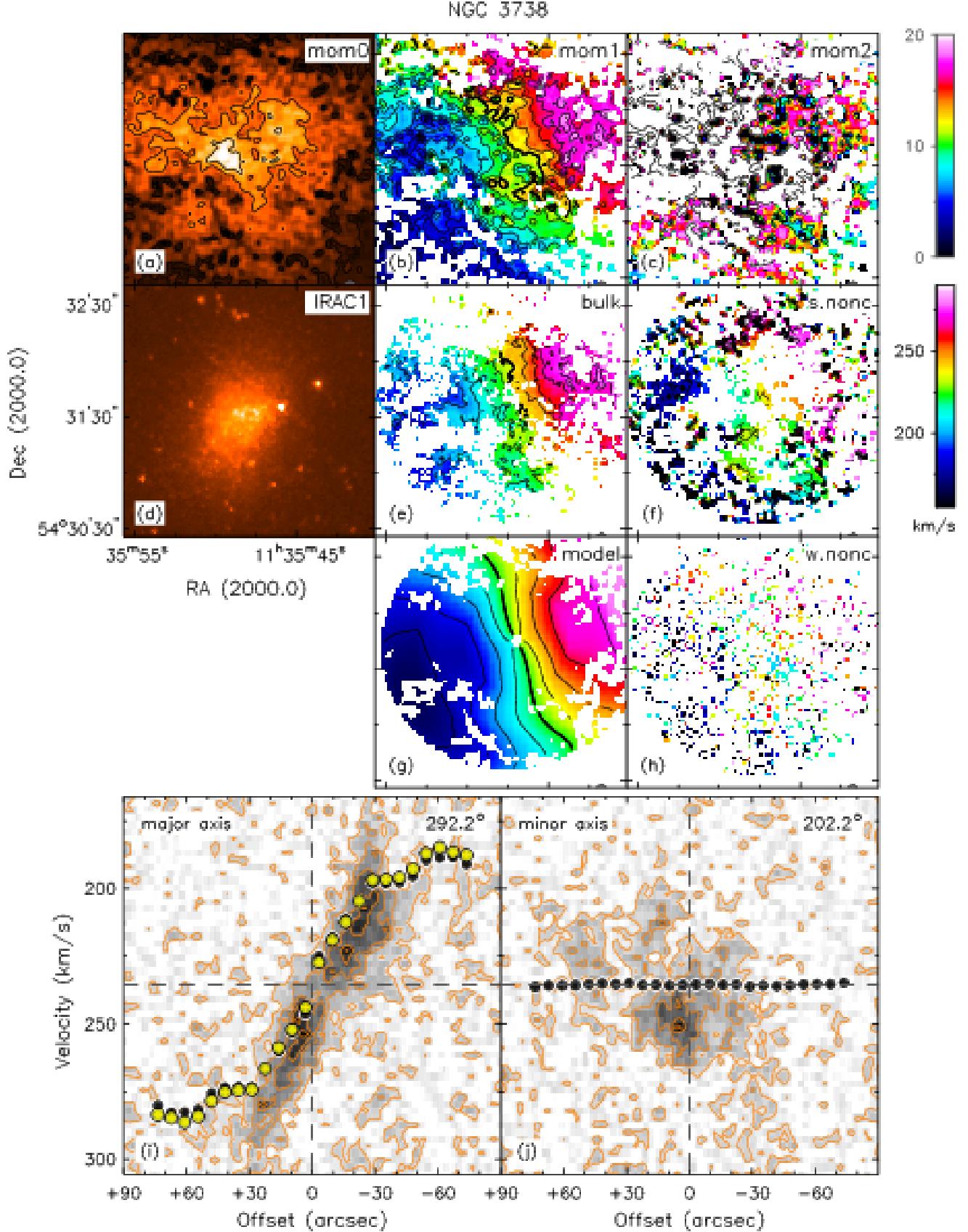


FIG. A.64.— HI data and *Spitzer* IRAC 3.6 μ m image of NGC 3738. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

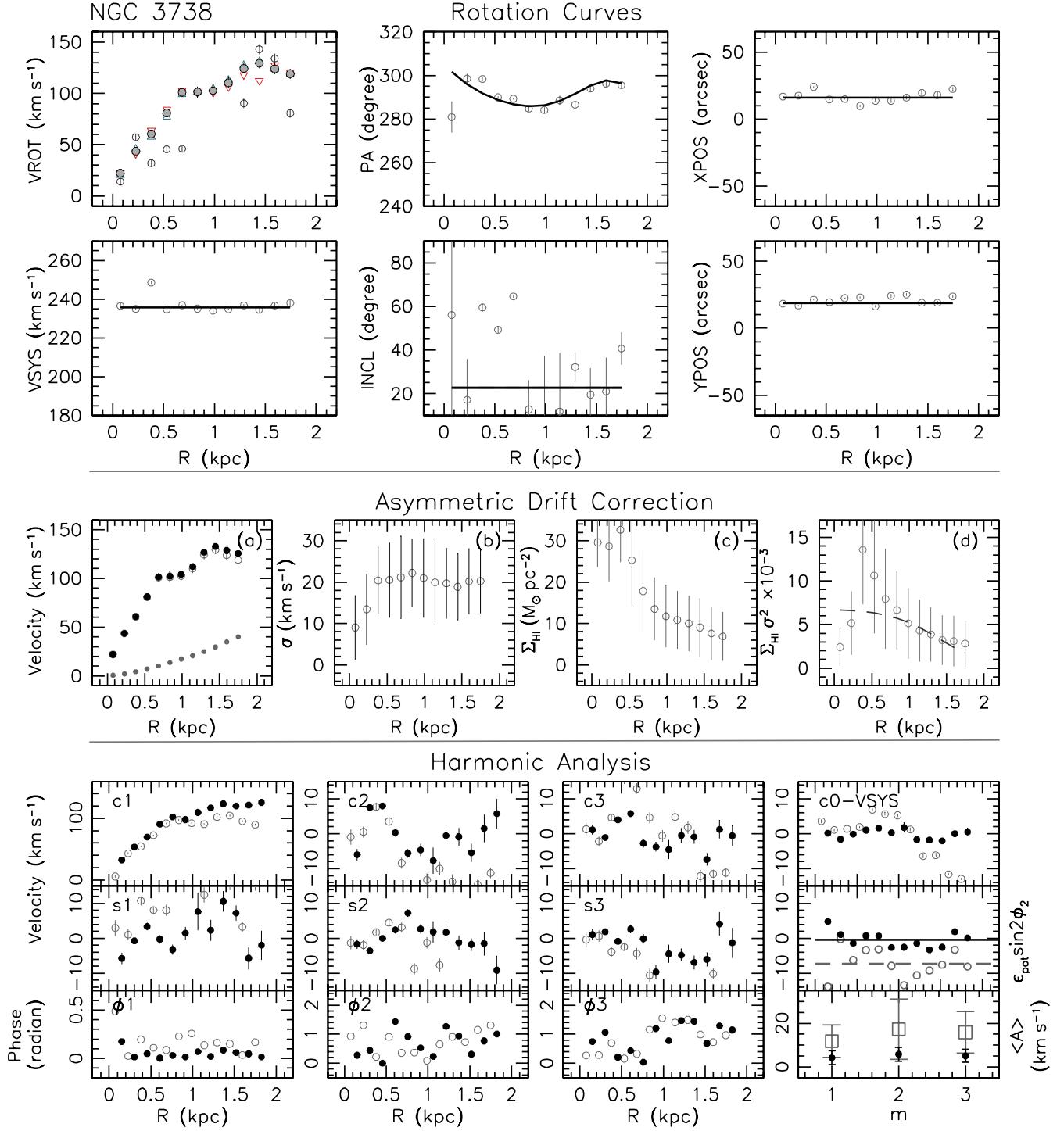
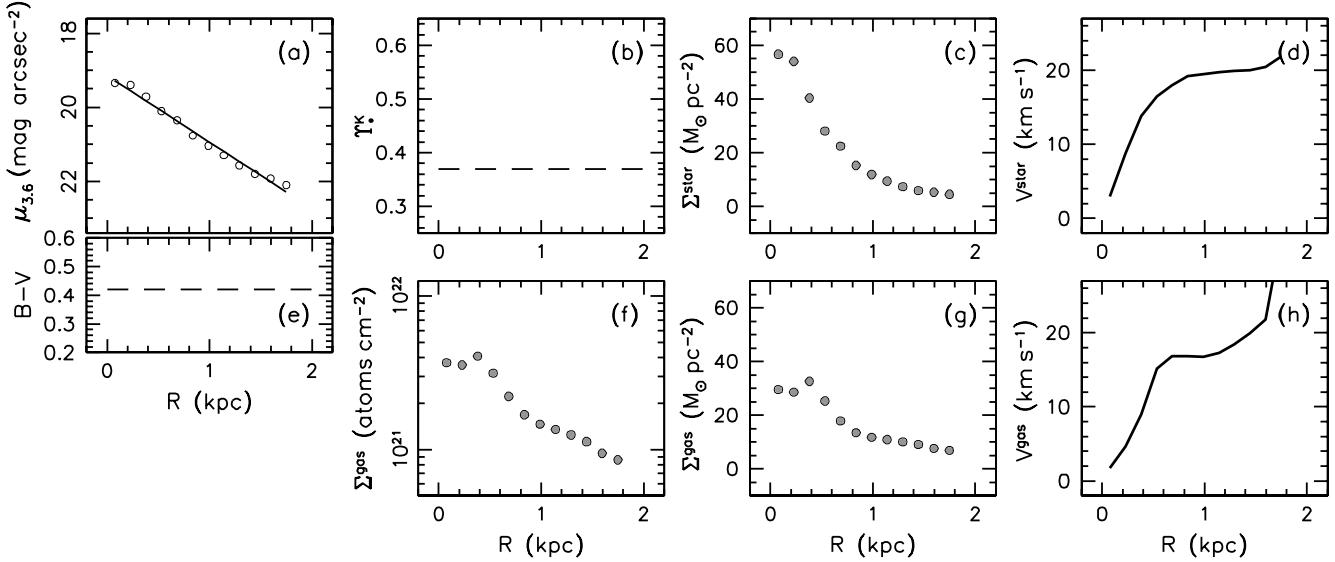


FIG. A.65.— Rotation curves, asymmetric drift correction and harmonic analysis of NGC 3738. See Appendix section B for details.

NGC 3738

Mass models of baryons



Disk–halo decomposition

Mass density profile

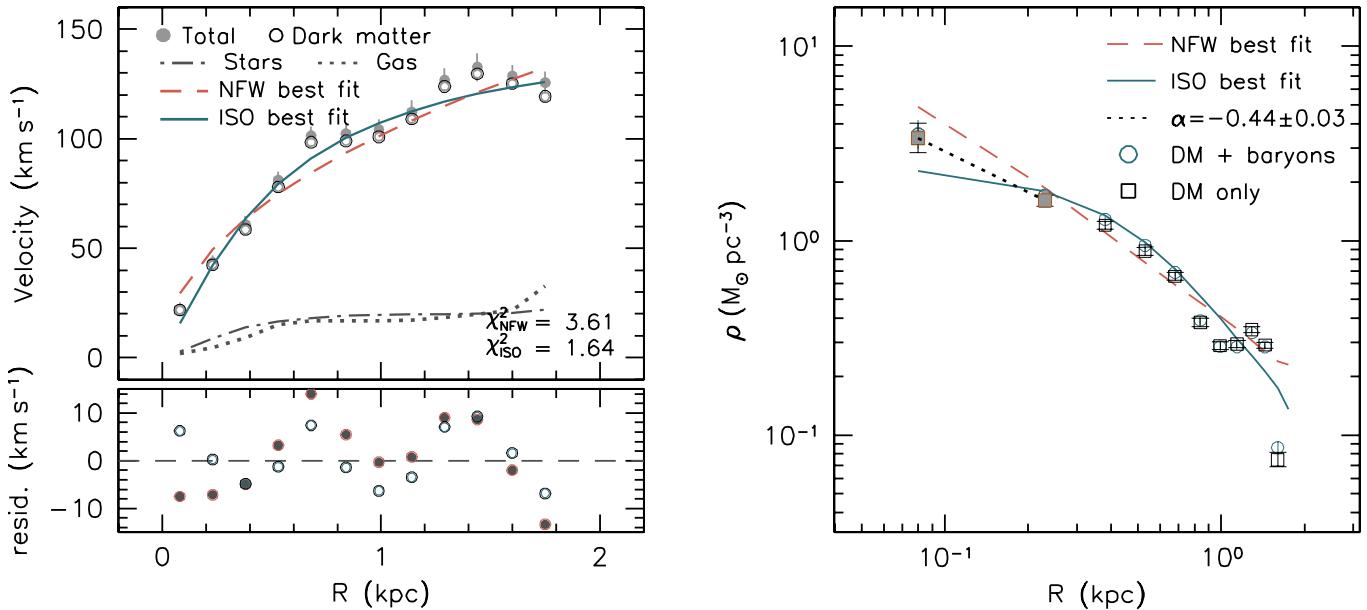


FIG. A.66.— The mass models of baryons, disk-halo decomposition and mass density profile of NGC 3738. Please refer to the text in Sections 3 and 4 for full information.

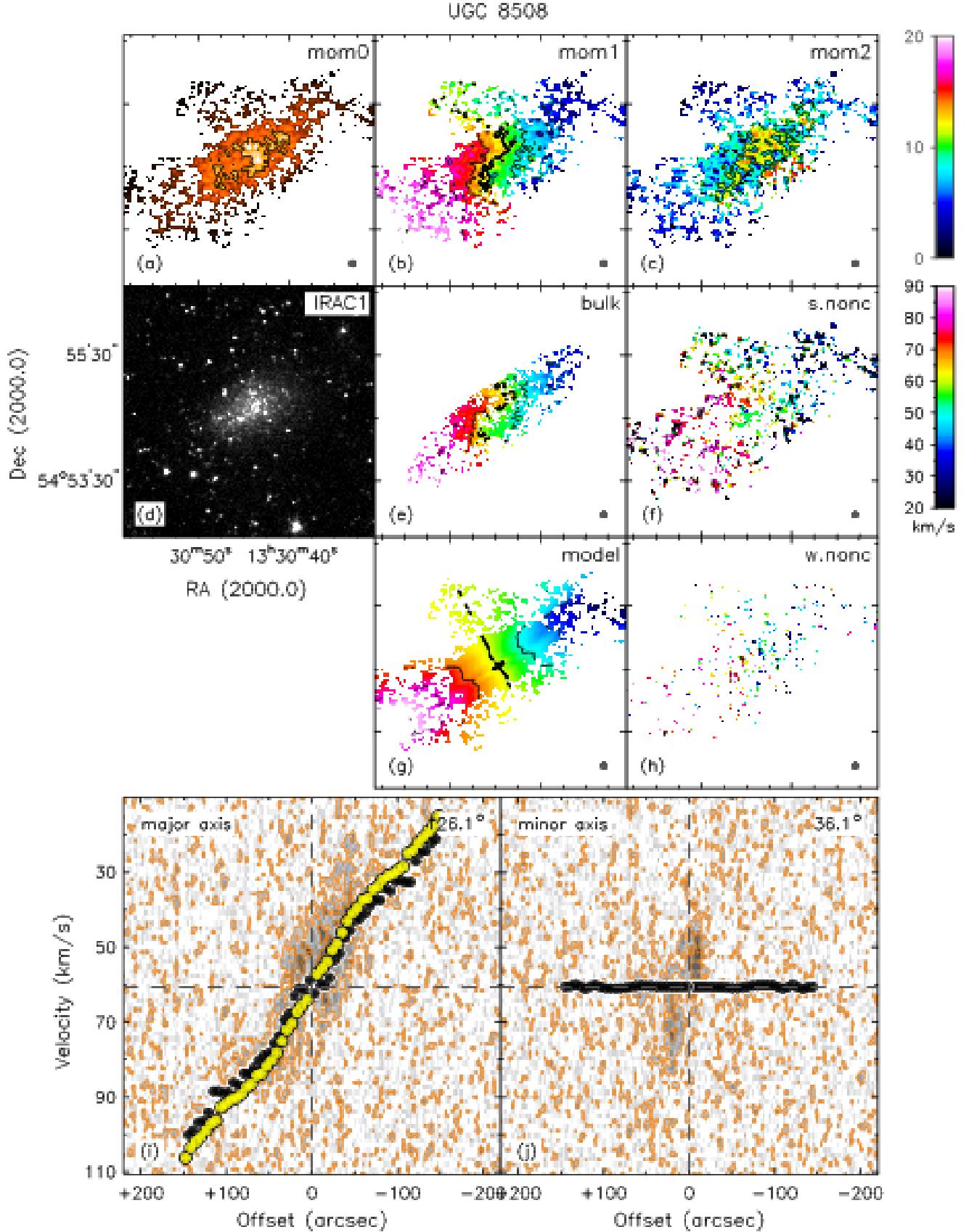


FIG. A.67.— HI data and *Spitzer* IRAC 3.6 μ m image of UGC 8508. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 10 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 10 km s $^{-1}$. See Appendix section A for details.

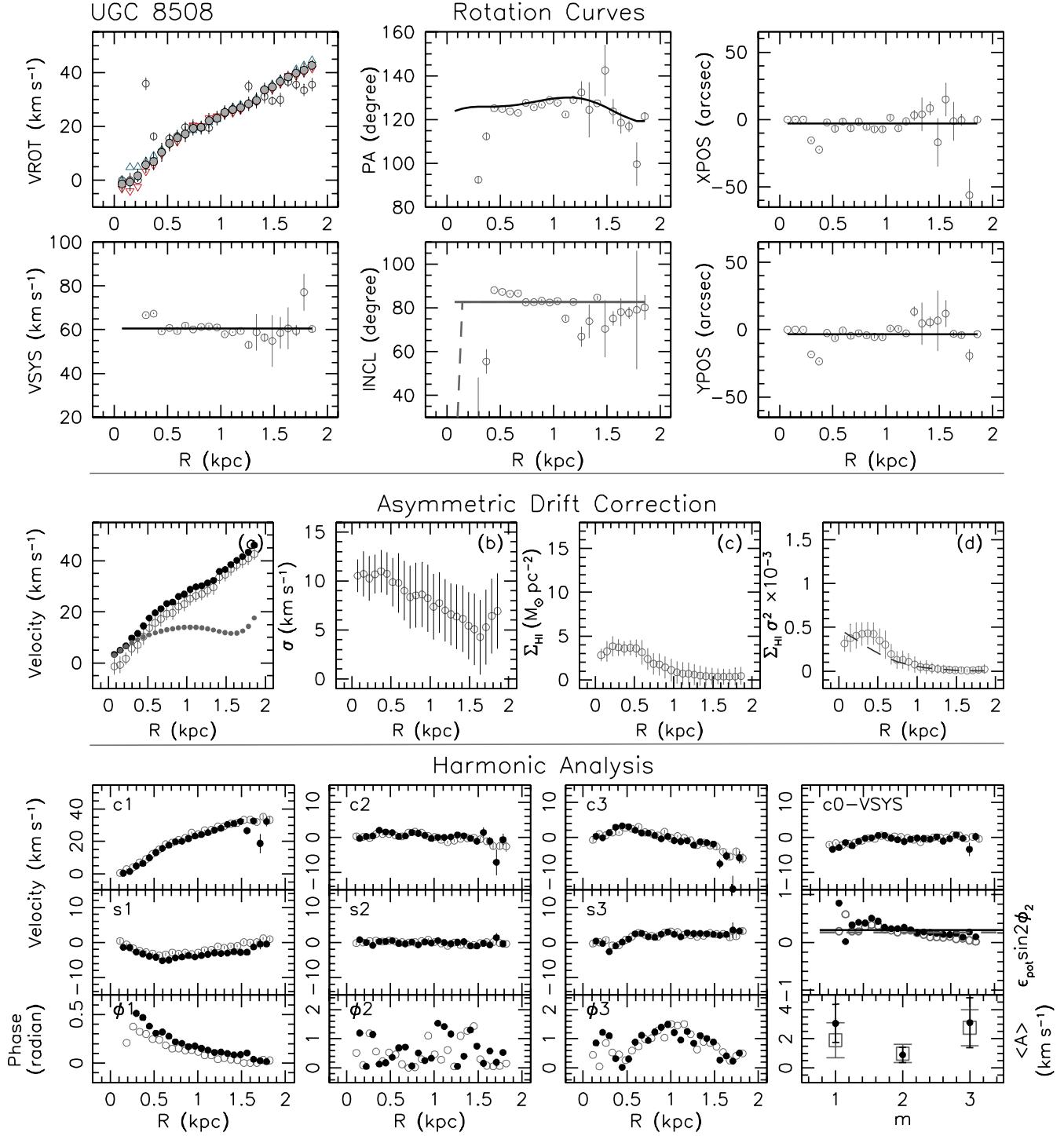


FIG. A.68.— Rotation curves, asymmetric drift correction and harmonic analysis of UGC 8508. See Appendix section B for details.

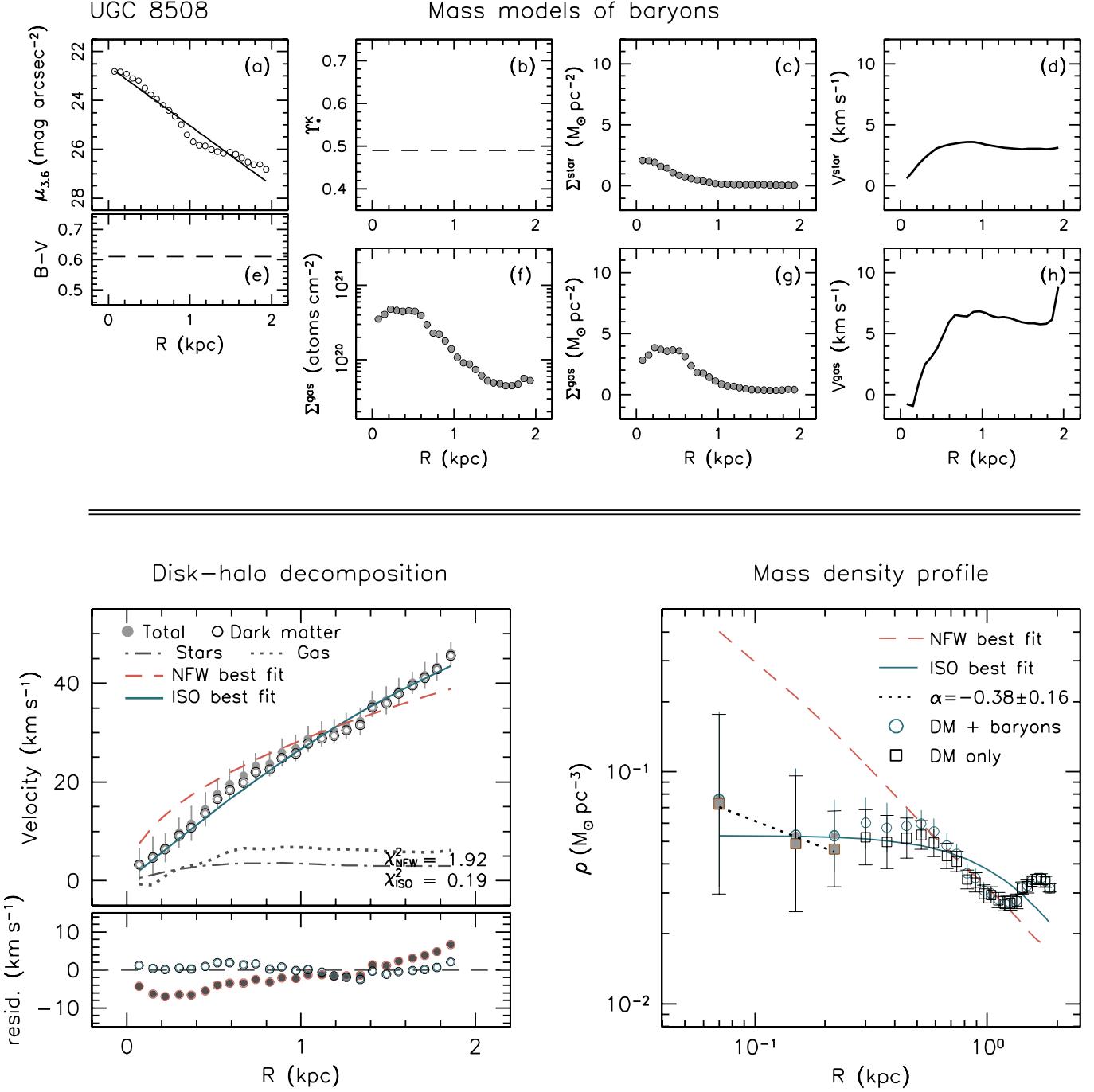


FIG. A.69.— The mass models of baryons, disk-halo decomposition and mass density profile of UGC 8508. Please refer to the text in Sections 3 and 4 for full information.

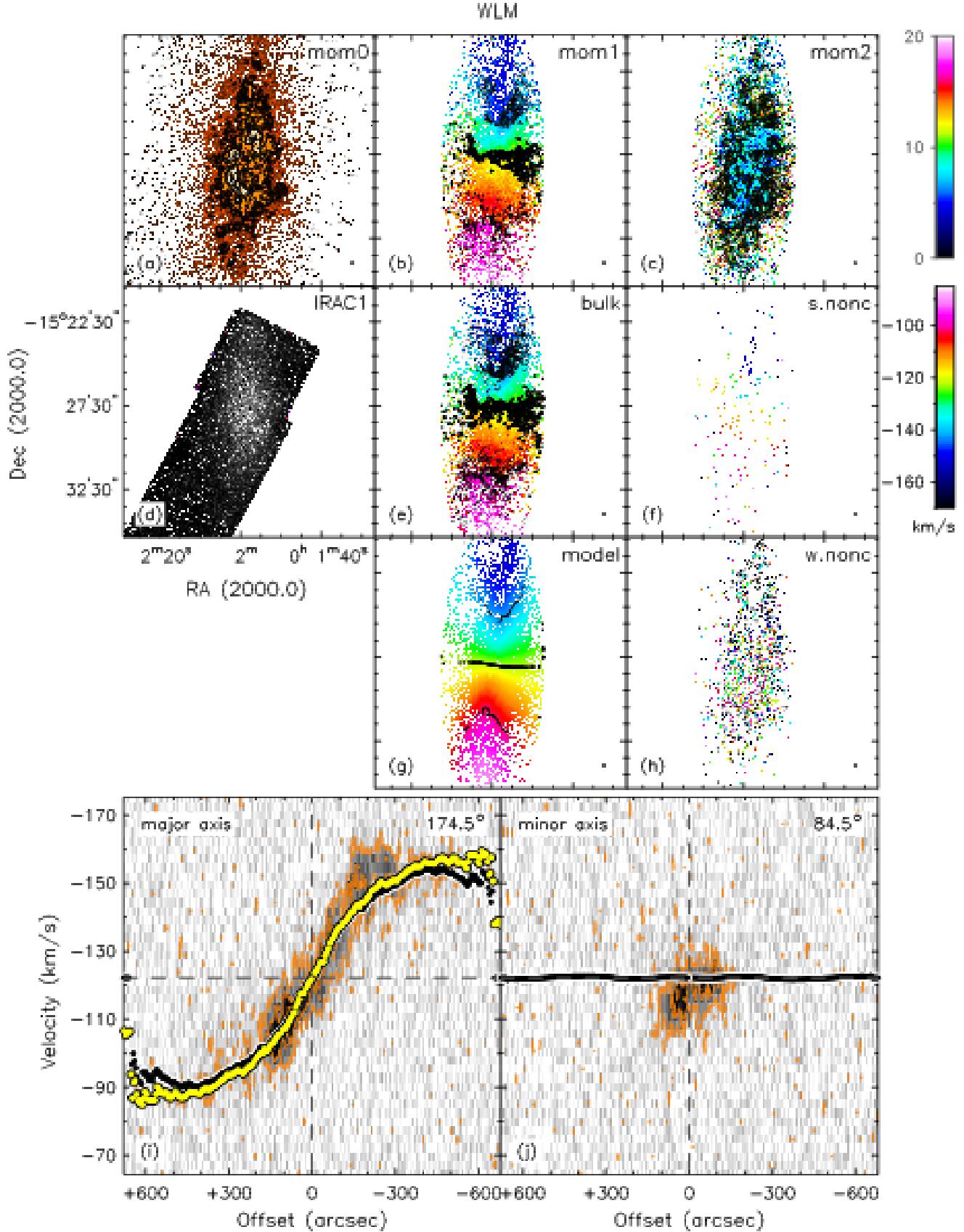


FIG. A.70.— HI data and *Spitzer* IRAC 3.6 μ m image of WLM. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 20 km s $^{-1}$. Velocity dispersion contours run from 0 to 20 km s $^{-1}$ with a spacing of 10 km s $^{-1}$. See Appendix section A for details.

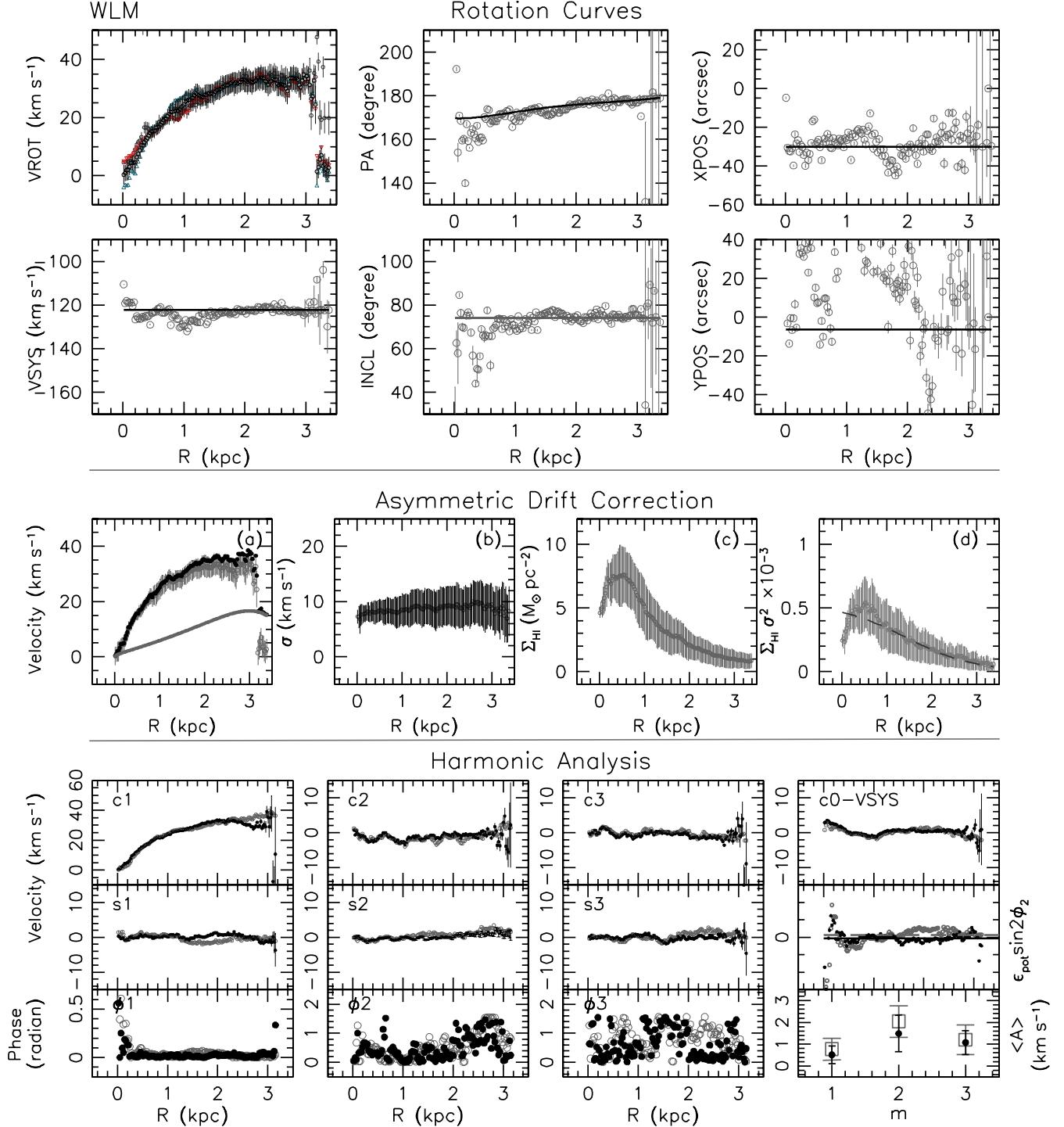


FIG. A.71.— Rotation curves, asymmetric drift correction and harmonic analysis of WLM. See Appendix section B for details.

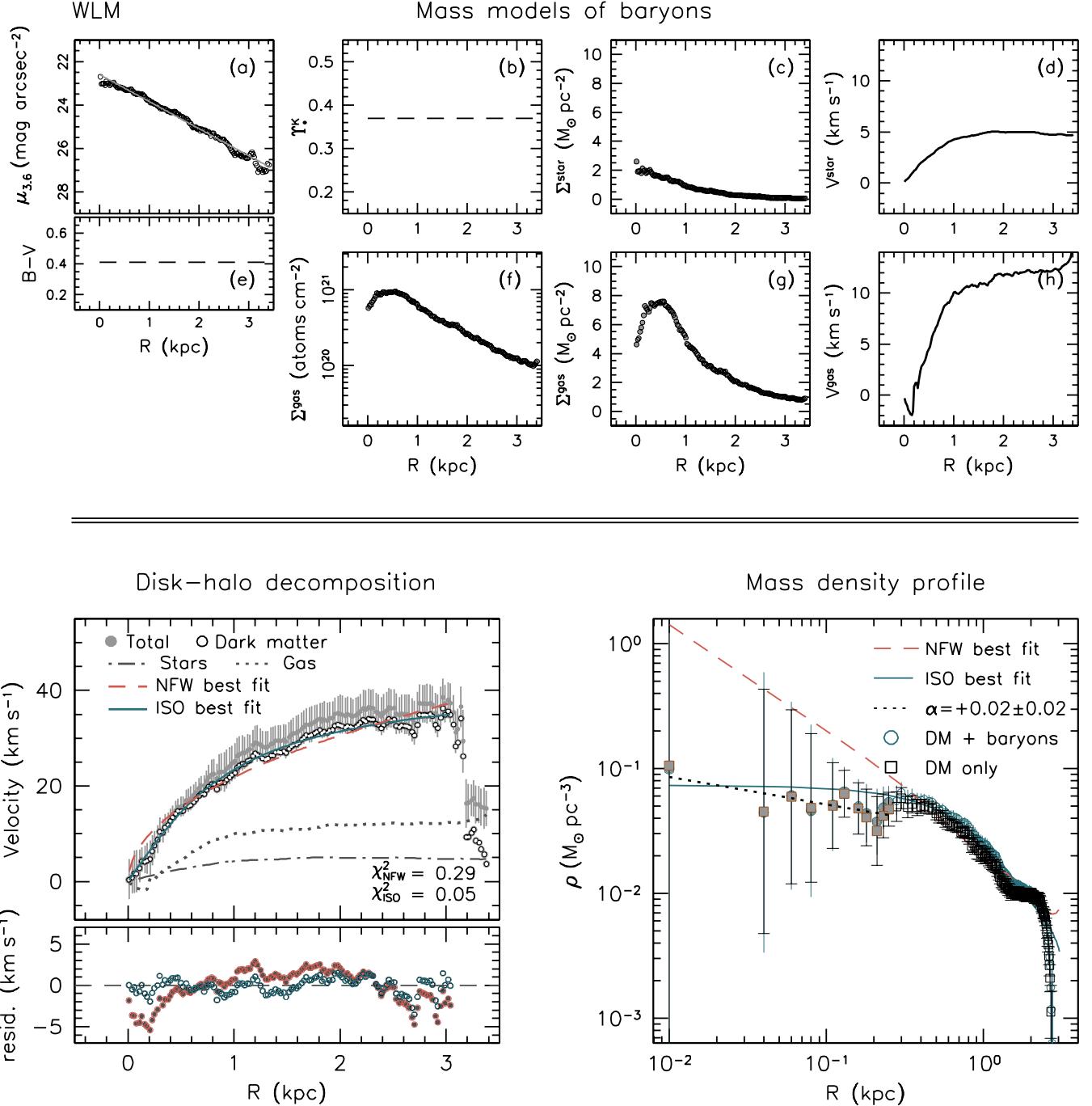


FIG. A.72.— The mass models of baryons, disk-halo decomposition and mass density profile of WLM. Please refer to the text in Sections 3 and 4 for full information.

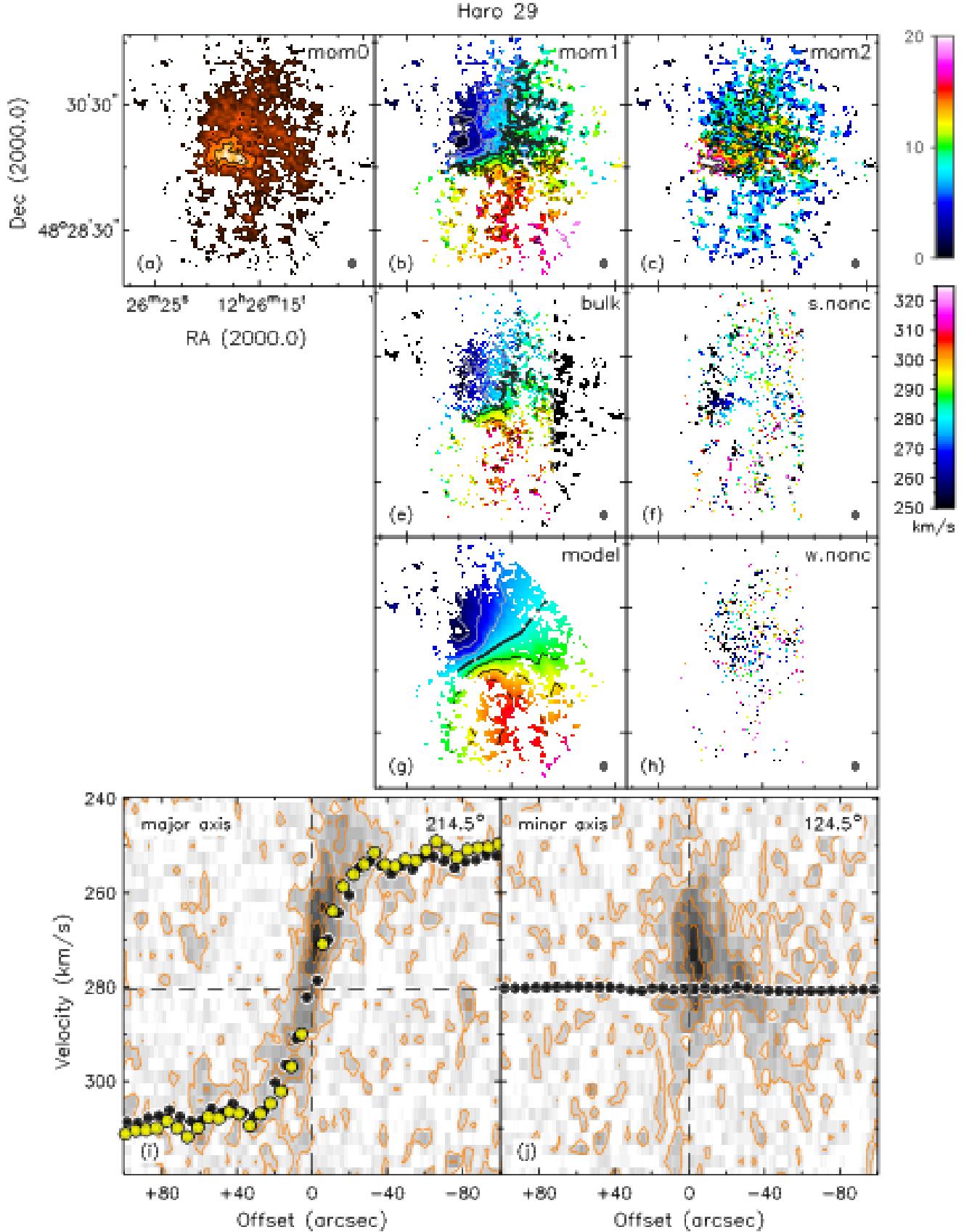


FIG. A.73.— HI data and *Spitzer* IRAC 3.6 μ m image of Haro 29. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s^{-1} . Velocity dispersion contours run from 0 to 30 km s^{-1} with a spacing of 10 km s^{-1} . See Appendix section A for details.

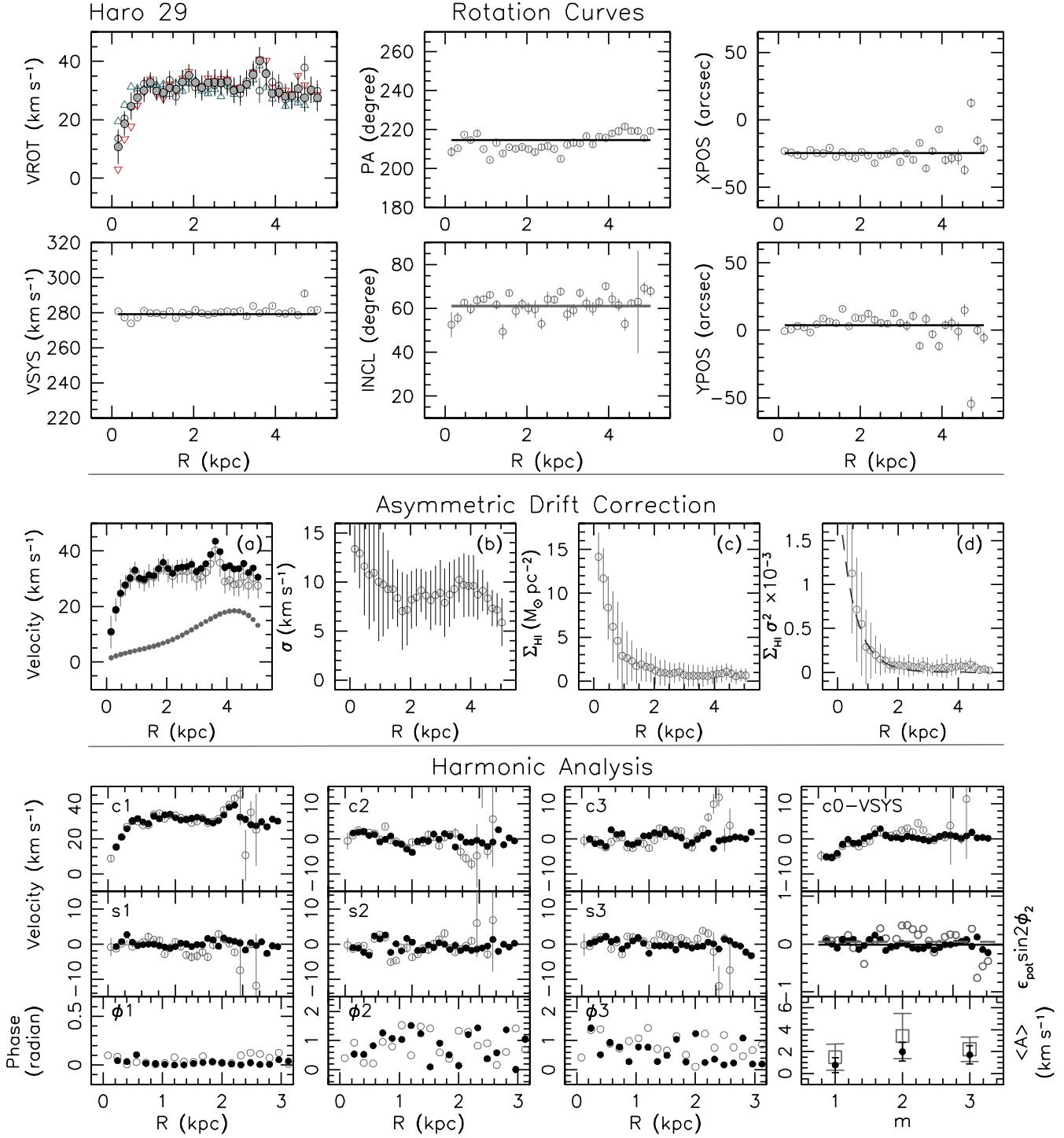


FIG. A.74.— Rotation curves, asymmetric drift correction and harmonic analysis of Haro 29. See Appendix section B for details.

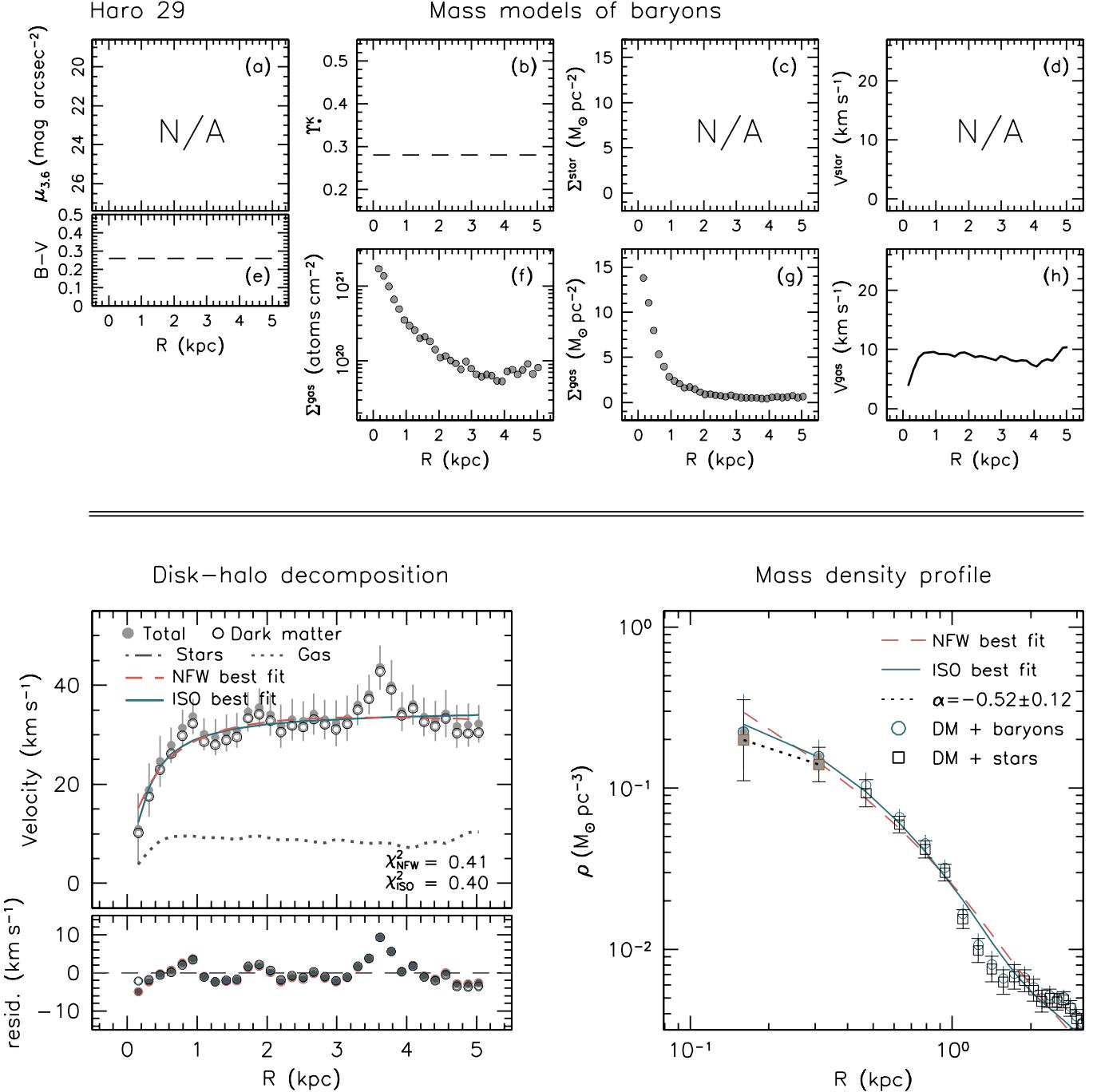


FIG. A.75.— The mass models of baryons, disk-halo decomposition and mass density profile of Haro 29. Please refer to the text in Sections 3 and 4 for full information.

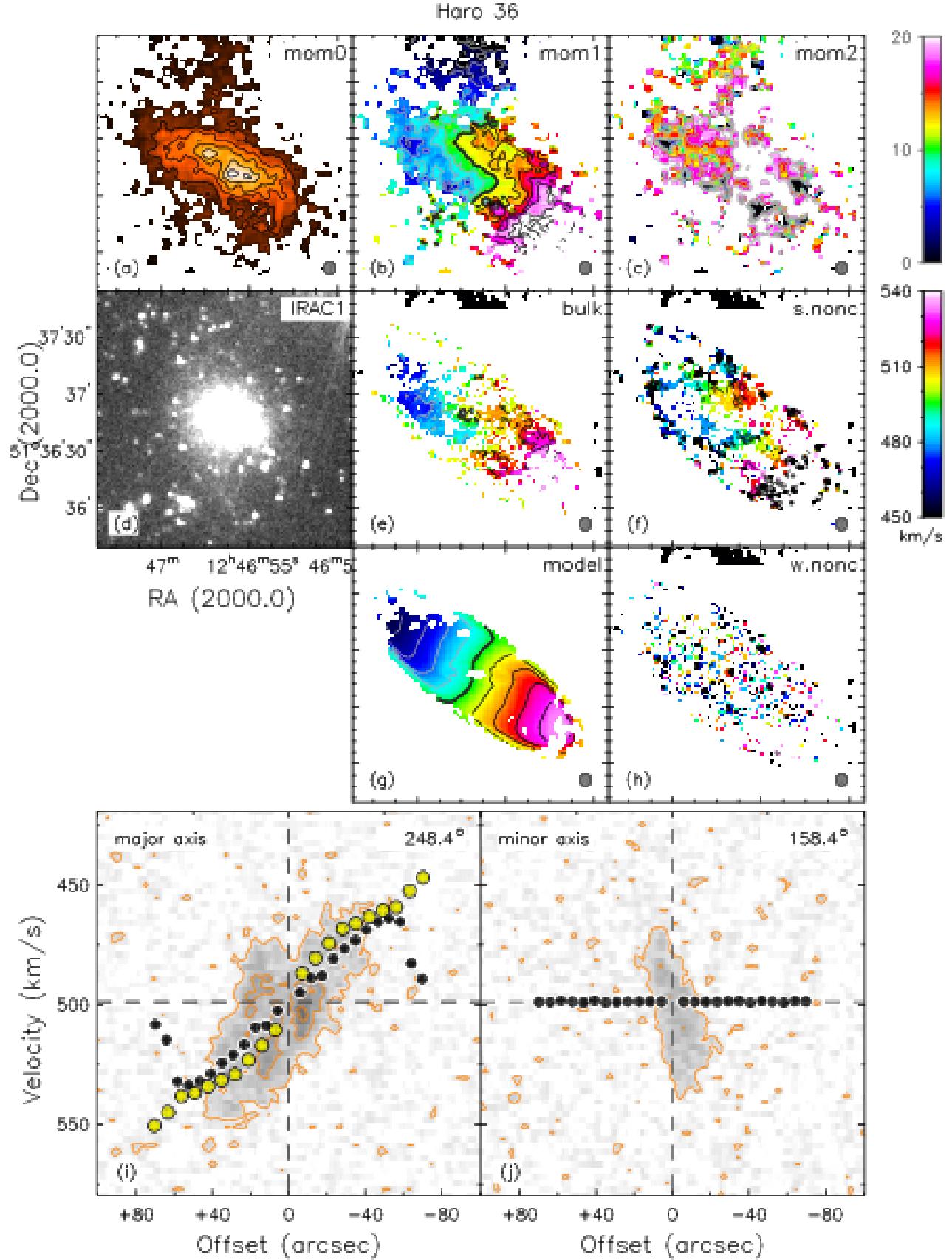


FIG. A.76.— HI data and *Spitzer* IRAC 3.6 μ m image of Haro 36. The systemic velocity is indicated by the thick contours in the velocity fields, and the iso-velocity contours are spaced by 8 km s^{-1} . Velocity dispersion contours run from 0 to 20 km s^{-1} with a spacing of 5 km s^{-1} . See Appendix section A for details.

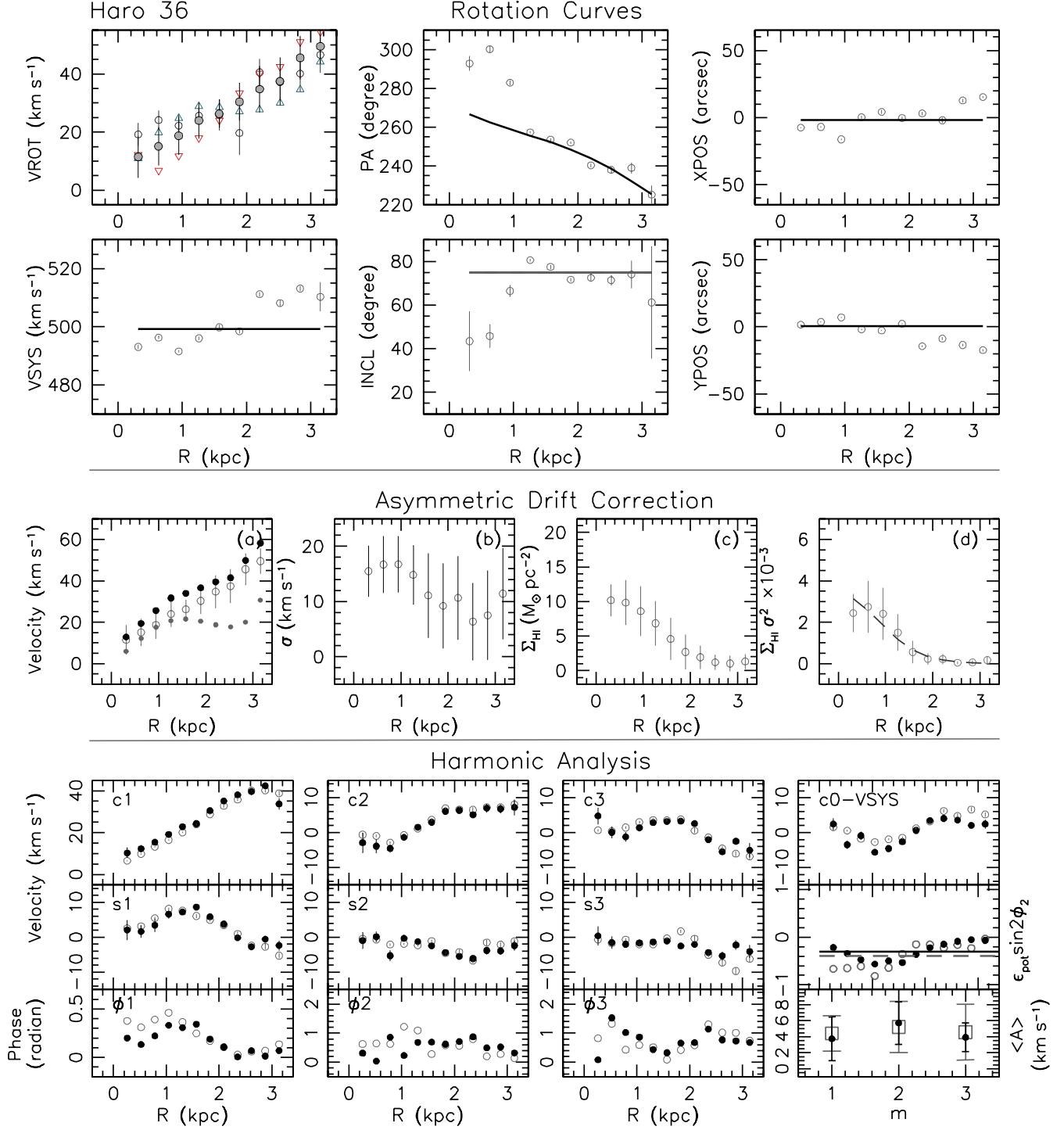


FIG. A.77.— Rotation curves, asymmetric drift correction and harmonic analysis of Haro 36. See Appendix section B for details.

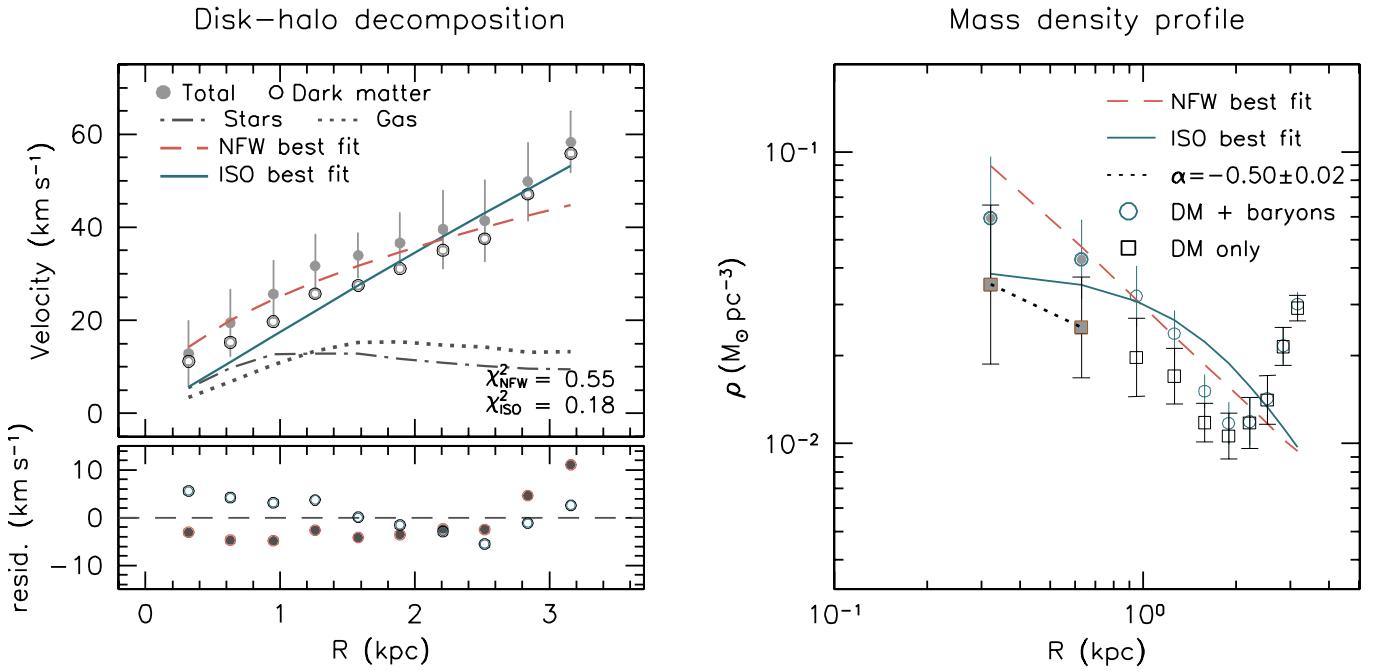
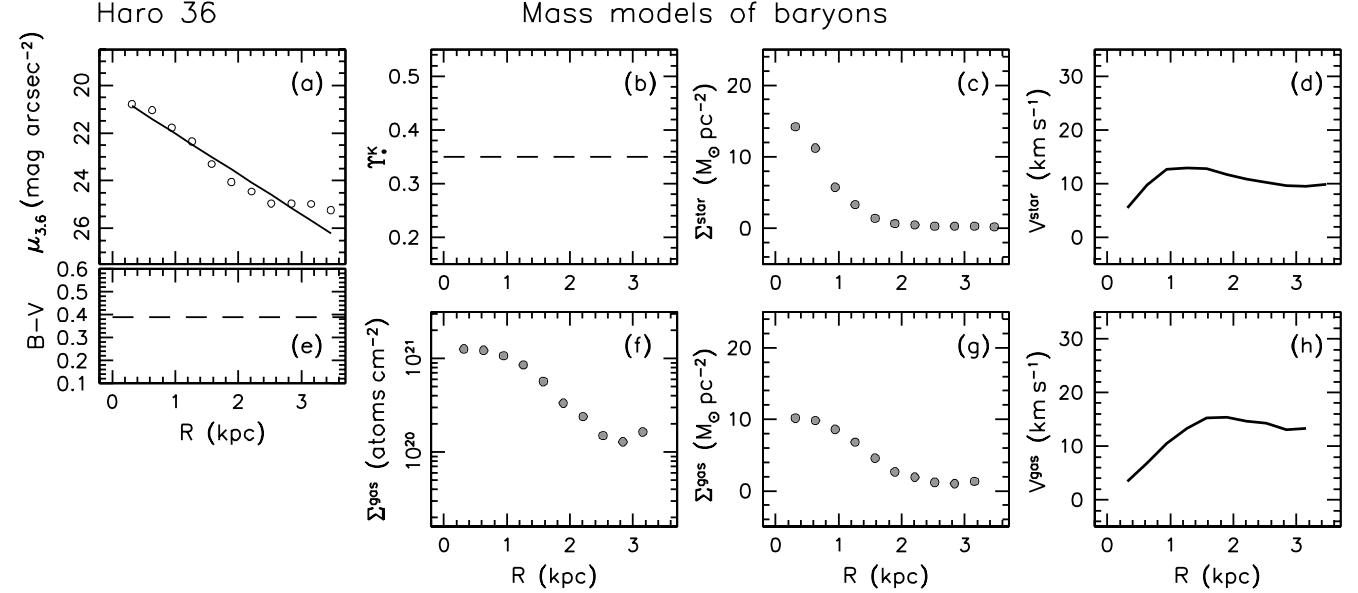


FIG. A.78.— The mass models of baryons, disk-halo decomposition and mass density profile of Haro 36. Please refer to the text in Sections 3 and 4 for full information.