

1. Science Justification

1.1. Cluster Mass Mapping

Science about cluster masses, substructure. Introduce CLASH.

1.2. Kinematic Weak Lensing

Weak lensing is weak, but we propose a novel technique to greatly increase the signal-to-noise using spectroscopic information. This approach was originally described by Blain (2002); Morales (2006).

Disk galaxies are inclined at random with respect to the line of sight to the observer, so the measured rotation velocity is related to the true circular velocity of the disk by $v_{\text{obs}} = v_{\text{circ}} \sin i$. Correcting for the effects of inclination has been an important observational complication inherent in TFR studies to date. Typically observers model the galaxy as a thin disk and estimate $\sin i$ from the ellipticity of the image.

After inclination correction, estimates for the intrinsic fractional scatter in v_{circ} at fixed luminosity or stellar mass are typically .05 dex or less. If the TFR is regarded as known, then a galaxy's velocity offset from the relation becomes an estimator for its true unlensed ellipticity. Properly weighted, the scatter in the ellipticity estimated this way is only .015 – in other words, knowledge of the disk inclination controls for 95% of the variation in the observed shape of the disk. This quantity is sensitive to measurement errors in v_{circ} , as shown in figure 1; in general, rotation curve errors of 10 km/s reduce the disk galaxy shape noise by an order of magnitude.

1.3. A Pilot Study for a Stage V Dark Energy Experiment

Weak gravitational lensing has been widely touted as a powerful probe of cosmology. It is a major science driver for several large imaging surveys, such as the Dark Energy Survey (DES) the Large Synoptic Survey Telescope (LSST), and space missions such as Euclid and the Wide-Field Infrared Survey Telescope (WFIRST). Weak lensing is weak, however, and all current measurements are limited by the scatter in galaxy shapes, which is an order of magnitude or more greater than the typical weak lensing signal. This means that weak lensing analyses must average over every available galaxy image, pushing the analysis to include faint and poorly-resolved galaxies for which unbiased measurements of galaxy shapes are difficult. Precision estimates of galaxy redshifts from photometry is crucial for these analyses, as the current generation of ongoing surveys will entail sample sizes of a few $\times 10^8$ galaxies, which is two orders of magnitude beyond the size of the largest existing spectroscopic samples. Percent-level biases in the photometric redshifts will have a large impact on the error budgets for this next generation of surveys.

Here we propose a pilot study for a new weak lensing measurement technique that uses the Tully-Fisher relation (TFR) to obviate the photo-z and shear calibration problems while offering a very large reduction in the effective lensing noise.

REFERENCES

Blain, A. W. 2002, ApJ, 570, L51

Morales, M. F. 2006, ApJ, 650, L21

2. Technical Remarks

2.1. Targets and Exposures

2.2. Backup Program

2.3. Supplementary Observations

2.4. Status of Previously Approved Keck Programs

2.5. Path to Science from Observations

2.6. Technical Concerns

2.7. Experience and Publications

2.8. Resources and Publication Timescale

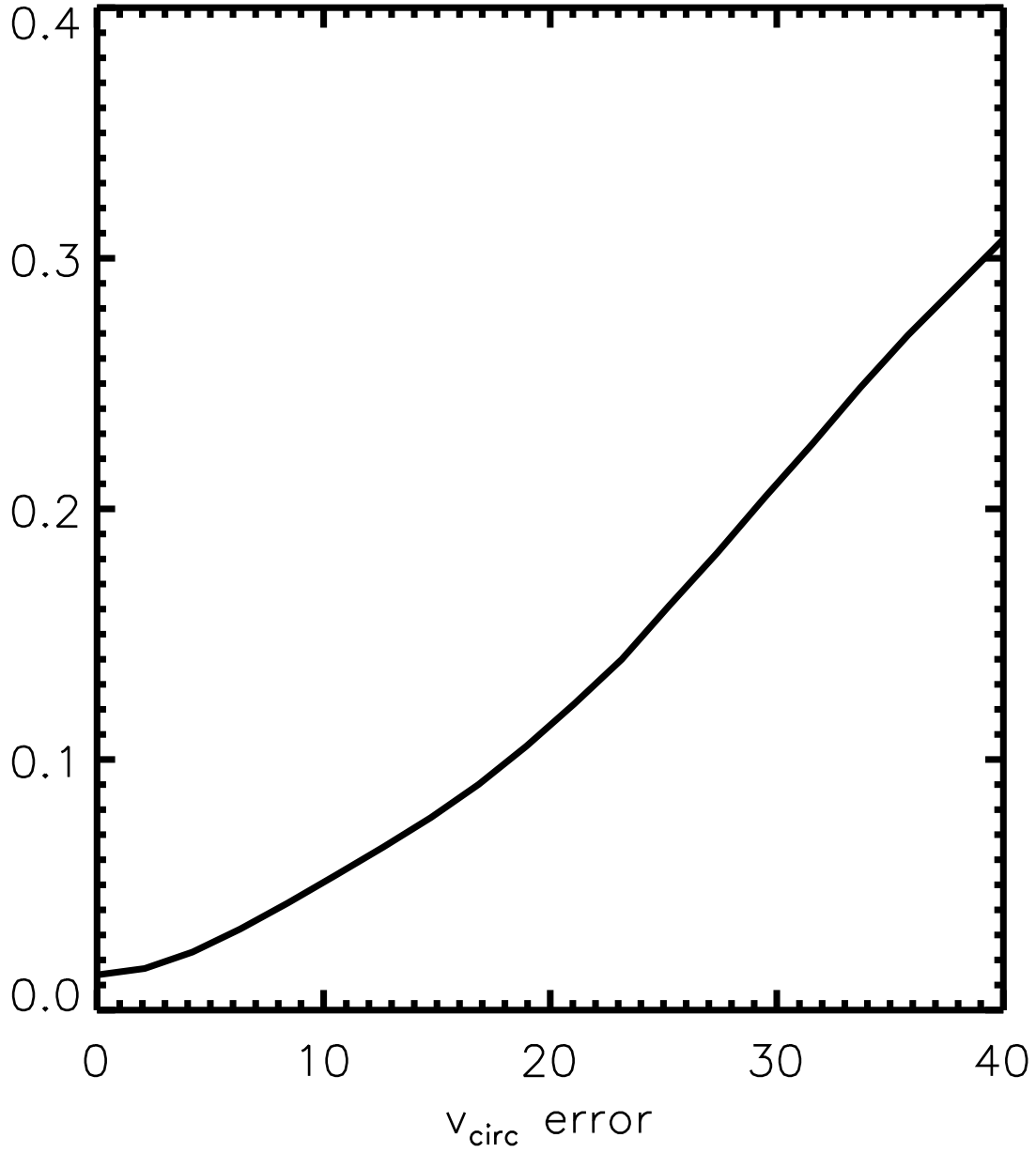


Fig. 1.— Effective shape noise σ_{TF} as a function of the measurement error on the disk circular velocity.