

The Photometric LSST Astronomical Time-series Classification Challenge (PLAsTiCC): Data set

We briefly describe the PLAsTiCC data set that will be tested by the Kaggle team. This note will be updated for the full release of the data to the community.

1. Introduction

While we think of the night sky as static, it is filled with sources of light that vary in brightness on timescales from seconds and minutes to months and years.

Some of these events are classified as *transients*, and can be the observational consequences of a large variety of astronomical phenomena. For example, the cataclysmic event that occurs when a supernova explodes generates a bright signal that fades with time, but does not repeat.

Other events are classified as *variables*, since they can vary their brightness in a periodic (or aperiodic) fashion, and originate from physical process governing high density regions of the Universe such as emission from the active galactic nuclei (AGN) at the hearts of galaxies, or our position in relation to a set of observed objects (e.g. eclipsing binaries).

These relatively rapid changing objects can provide important clues about themselves and their environment - as well as the evolution of the universe as a whole (e.g. type Ia supernovae provided the first evidence of the current accelerated expansion of the Universe caused by dark energy). Crucially, the different types of transients and variables, given *different* pieces of information, and therefore, the proper classification of transients is a crucial task in observational astronomy - specially in the light of large data volumes expected for the next generation of astronomical surveys.

The main aim of this challenge is: *can one classify astronomical transients and variables from a photometric data set designed to mimic the data from the upcoming Large Synoptic Survey Telescope (LSST)?* Crucially, the classification will occur on a large test set, but the training data will be a small, and poorly representative training set, to mimic the challenges we face observationally.

We now give more detail on LSST, and the challenge at hand. The two modes for characterising the light from the objects, are called spectroscopy and photometry.

Spectroscopy is the modern equivalent of using a prism to separate a beam of white light in its composite (e.g. rainbow) colours. It is a high resolution measurement which allows us to identify any emission/absorption features indicative of specific chemical elements present in the object - and are also the main source of information enabling classification. Despite being paramount for the classification task, spectroscopy is an extremely time consuming process - with integration times hanging from 20 minutes to a few hours depending on the telescope and brightness of the source.

Traditionally, classification is completely based on spectroscopic measurements - however, given the volume of data expected from the upcoming large scale sky surveys, this strategy is not sustainable. An alternative approach, based on cheaper/lower resolution measurements must be found. The spectra are described as flux per wavelength interval.

Different physical processes (and chemical elements) are found at different wavelengths of the spectrum, and so different parts are useful to distinguish between objects.

Photometry gives the overall information of how bright the source is (how much energy the object is emitting) at a particular moment. It can be described as the low resolution counterpart of spectroscopy - or the integral of all the information stored in the spectra. The resultant photometric magnitudes m are related to the flux in a given (average) wavelength through

$$F = 10^{-0.4m+11} + \text{noise} \quad (1)$$

These band fluxes are the integrals of the appropriately redshifted spectrum over the filter bandpasses of atmosphere and the instrument divided by the energy of photons in the central wavelength of the filter. A sequence of photometric observations made at different times is called a light curve. It holds the information of how the energy of the source

evolves with time and can also be used to characterize different types of supernovae (although the distinction between classes is much more discrete than we would hope it to be). As a consequence, for each object we will have a number of light curves in each filter (or band). For LSST, the filters are the *ugrizy* filters and are roughly 1000 Angstrom wide, spanning a range from 3000 Ang to 9000 Ang. The observations are noisy, mostly because of sky noise which is due to light from different components, and this varies with time (eg. twilight, effect of the moon) and its strength is different in different bands. Therefore the bands will vary in usefulness for different objects. In addition, the bands are sensitive to different physical processes in the same way as the spectrum, and so will be more or less useful to distinguish between objects.

In addition to being more cost effective, it is possible to obtain photometry of objects to greater distances (and volumes) in the universe. For an object of given brightness (luminosity), the flux received on earth decreases with the distance to the object as

$$F = \frac{L}{4\pi d_L^2}. \quad (2)$$

New instruments, like the Large Synoptic Survey Telescope (LSST) - scheduled to begin observations in 2022 - will produce an unprecedented number of light curves. Once these transients are detected, we rely on agreements with other telescopes in order to acquire a small number of spectroscopic observations.

The final connection to a model of the cosmology of the universe, and the distances to objects comes through the redshift, z . Redshift is an empirical quantity that is defined by measuring the difference in the observed wavelength λ_o of a given feature (e.g. in the spectrum described above) compared to the emitted wavelength λ_e , or

$$z = \frac{\lambda_o - \lambda_e}{\lambda_e}. \quad (3)$$

It is the light analogy of the Doppler effect that we are used to in sound. It is clear then that using a spectrum to determine the redshift of an object gives the most precise result (with the smallest error σ). However photometry can also be used to determine a redshift for an object, the so-called photometric redshift, with a larger uncertainty. Given that photometric error relies on prominent features moving between the (large) bandpasses,

it is also possible to insert ‘catastrophic’ redshift error, where the redshift of the object is misassigned. These errors are rare (roughly 2% of the total number of objects), however can pose serious problems for classification of objects.

The distance modulus is connected to the cosmology as a function of redshift z and the components of the universe. We do not expect participants of this challenge to compute these functions, and will provide the distance modulus and error on the modulus as if the data had been run through a light-curve model fit.

We will describe the data in the following sections, and discuss the metrics used to classify objects in a separate note.

2. The data

The lightcurve data is made of non-homogeneously sampled, non-periodic time series with correlated errors obtained in several wavelength filters.

A database over all objects will be provided. The following data are provided:

- `objid`: the Object ID
- `ra`, `dec`: right ascension and declination (the sky coordinates of the object) (???)
- `newPtrmin`: the minimum Petrosian magnitude: an indication of sky noise
- `newPtrmax`: the maximum Petrosian magnitude: an indication of sky brightness
- `mwebv`, `mwebverr`: the value of the Milky Way extinction at that position on the sky, and its associated error
- `specz`: the spectroscopic redshift
- `photoz`, `photozerr`: the photometric redshift and its respective redshift error (these entries are null for the training set)

It will also contain a table of photometry with the following information:

- `mjd`: the Modified Julian Date (MJD) of the observations
- `pb`: the Petrosian
- `flux`: the measured flux in each of the LSST bands
- `fluxerr`: the error on the measured flux in each of the LSST bands
- `photflag`: photometry flag, which indicates the quality of the object (eg. it will be flagged if observations have saturated)

The training set will also contain the simulated model name, while the test data will not: it is the task of those participating to determine the classification flag of the test set.

2.1. Training data

The training data will be a subset of ~ 5000 objects taken from the larger test data set. The relative rates of the data will be adjusted slightly in the training sample. In addition, the test data will contain some rare objects that are not seen in the training data. The final classification classes will therefore contain an 'other' class to be used on any object that the classifiers have not seen before.

3. Conclusion

Acknowledgments

The DESC acknowledges ongoing support from the Institut National de Physique Nucléaire et de Physique des Particules in France; the Science & Technology Facilities Council in the United Kingdom; and the Department of Energy, the National Science Foundation, and the LSST Corporation in the United States. DESC uses resources of the IN2P3 Computing Center (CC-IN2P3—Lyon/Villeurbanne - France) funded by the Centre National de la Recherche Scientifique; the National Energy Research Scientific Computing Center, a DOE Office of Science User Facility supported by the Office of Science of

the U.S. Department of Energy under Contract No. DE-AC02-05CH11231; STFC DiRAC HPC Facilities, funded by UK BIS National E-infrastructure capital grants; and the UK particle physics grid, supported by the GridPP Collaboration. This work was performed in part under DOE Contract DE-AC02-76SF00515.