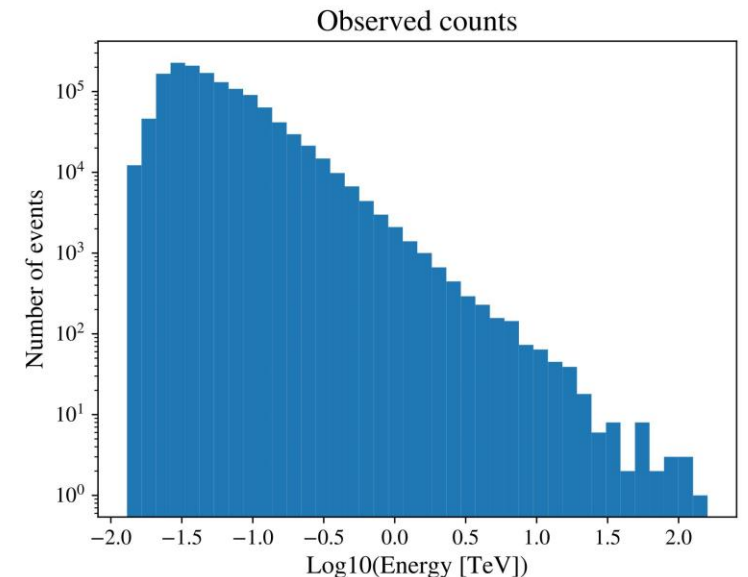
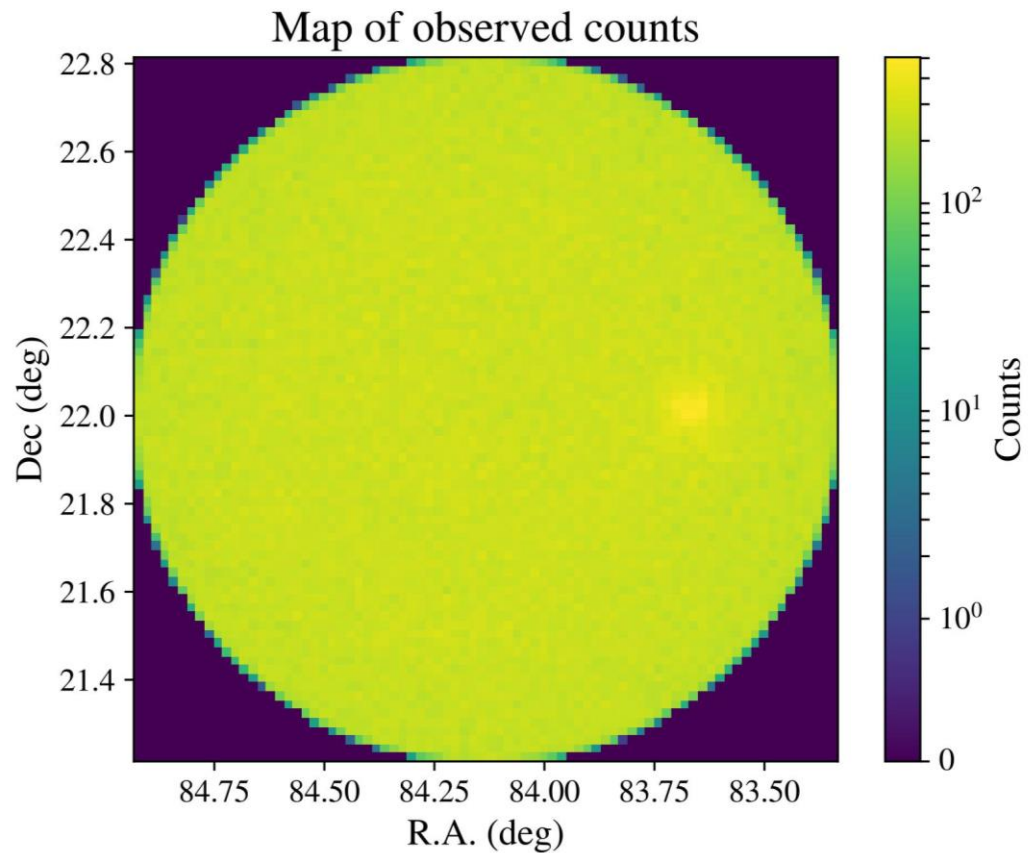


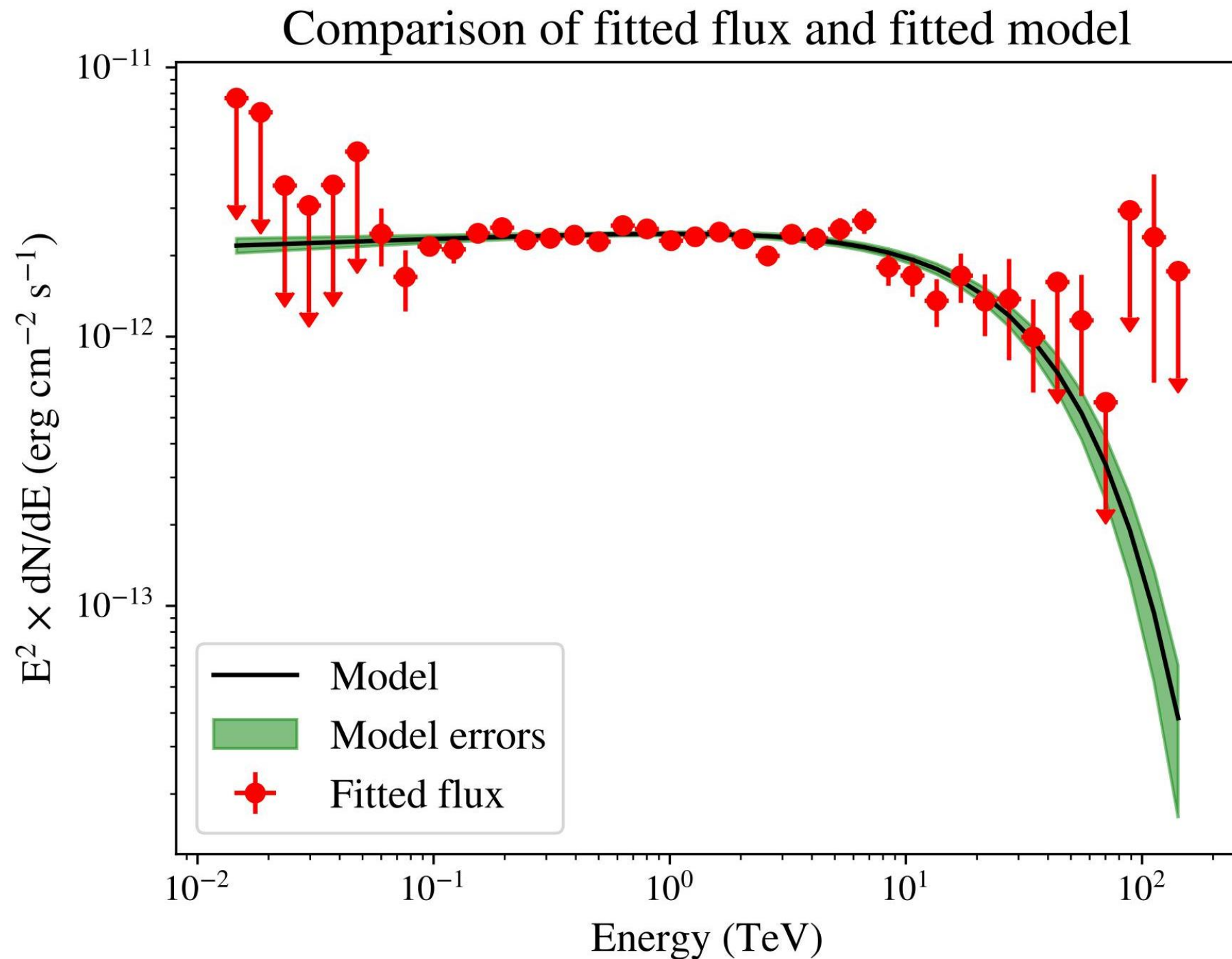
Intro: Simulations

- Point-like
- CTA IRF background
- 40h
- 40mCrab
- 50 TeV cutoff
- 1.6deg diameter FoV
- 0.5deg offset
- -2 index
- Edisp: on



Intro: Fitting

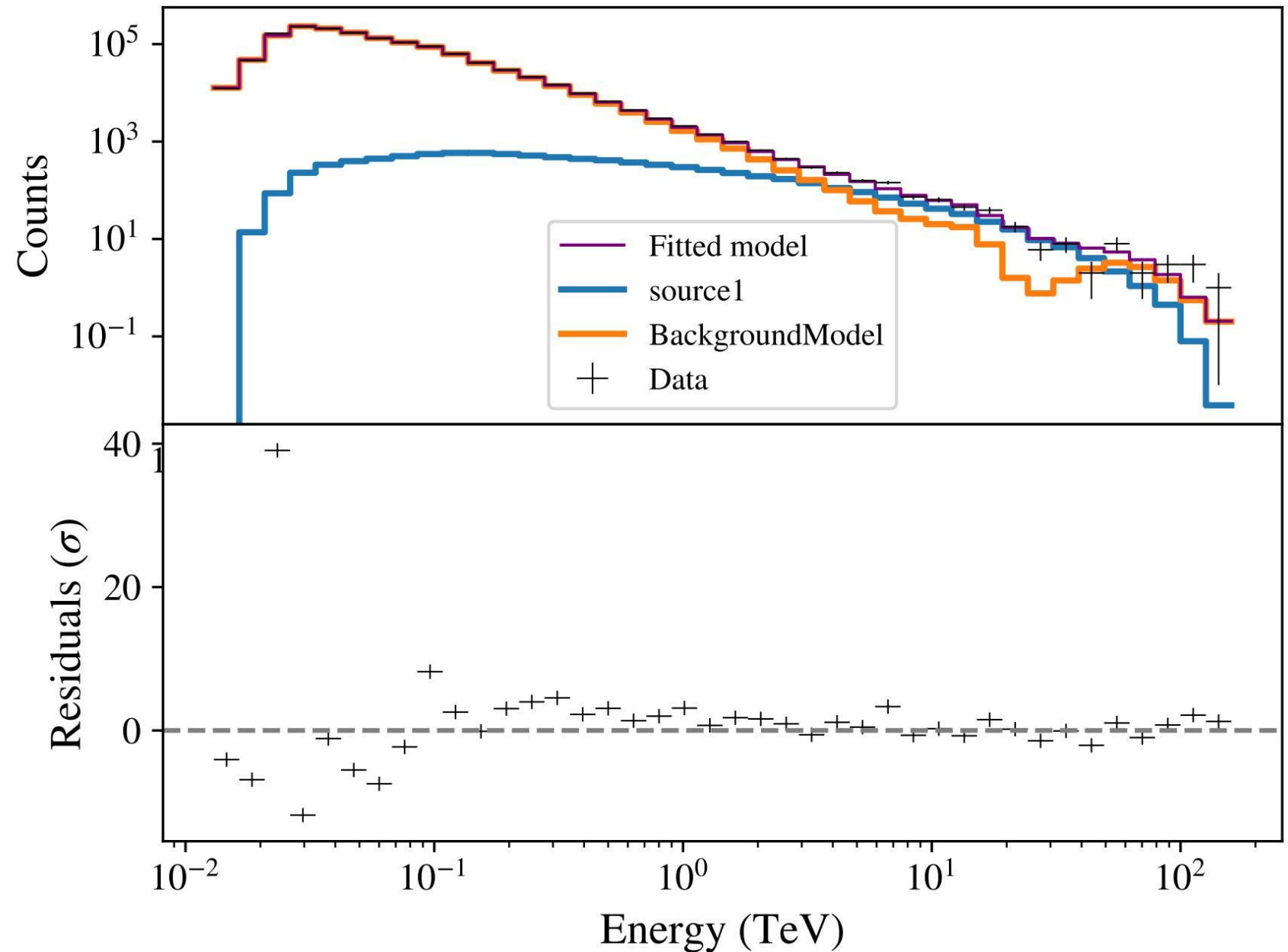
- 10 bins / energy decade
- 0.02 deg spatial resolution
- Fit by ctools' "ctlike"



Intro: Fitting ct'd

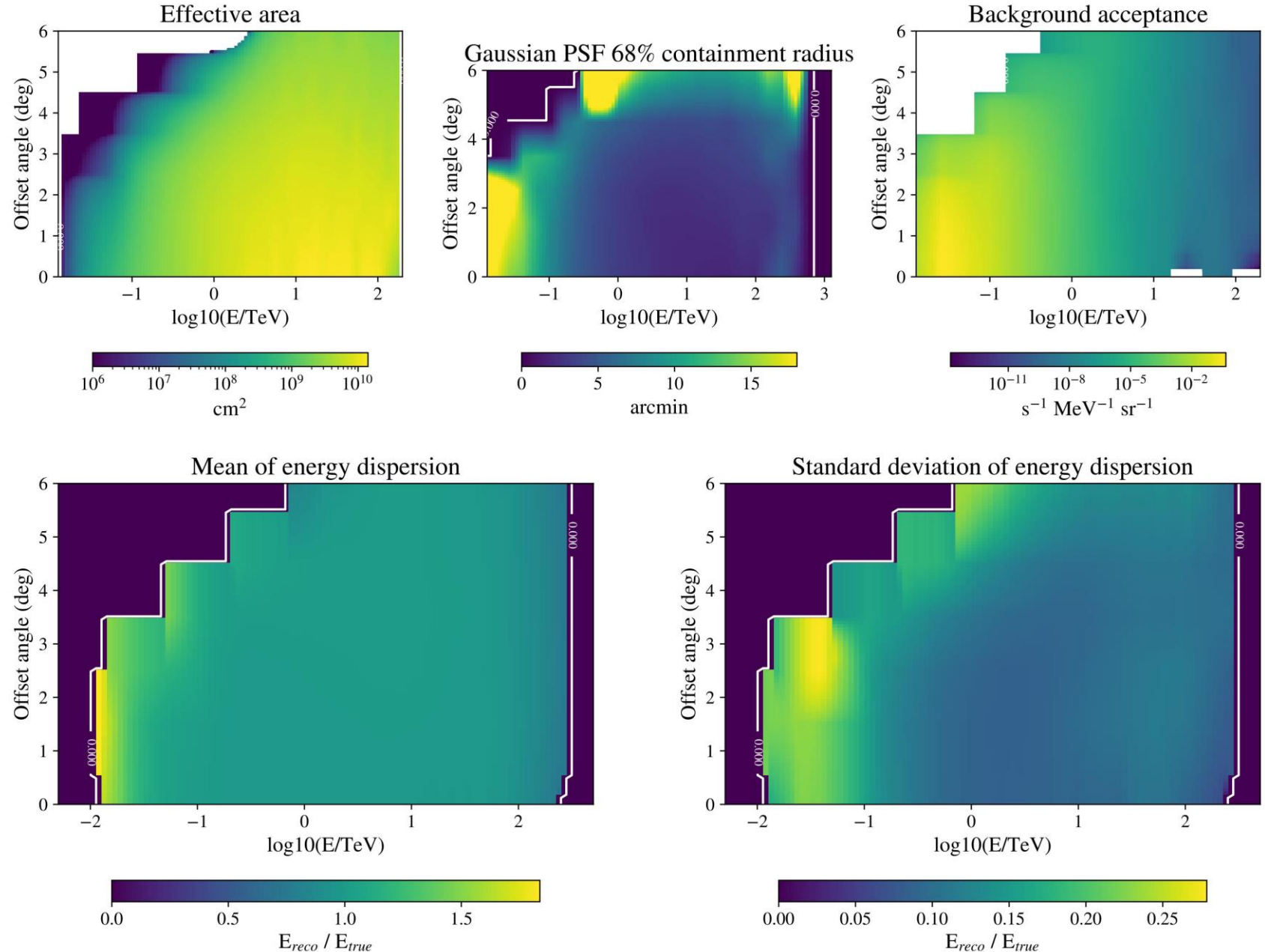
- 10 bins /
energy
decade
- 0.02 deg
spatial
resolution
- Fit by
ctools'
"ctlike"

Counts: observed and fitted model



Bracketing IRFs

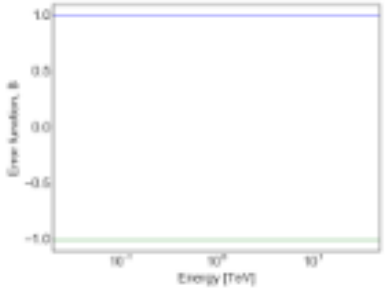
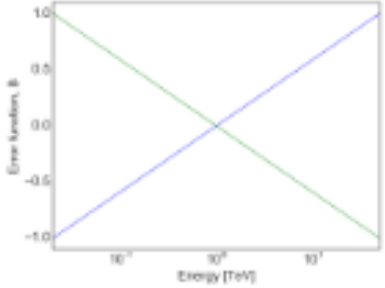
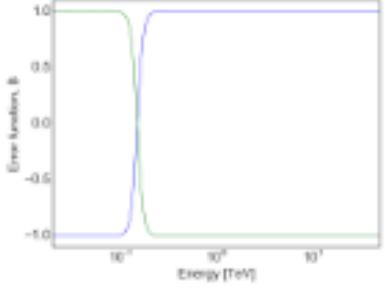
- PSF is not bracketed
- Only scaling, no resolution changes
- AEff 5% scaling
- EDisp 6% scaling
- The counts, spectra, models come out different
- Original IRF →

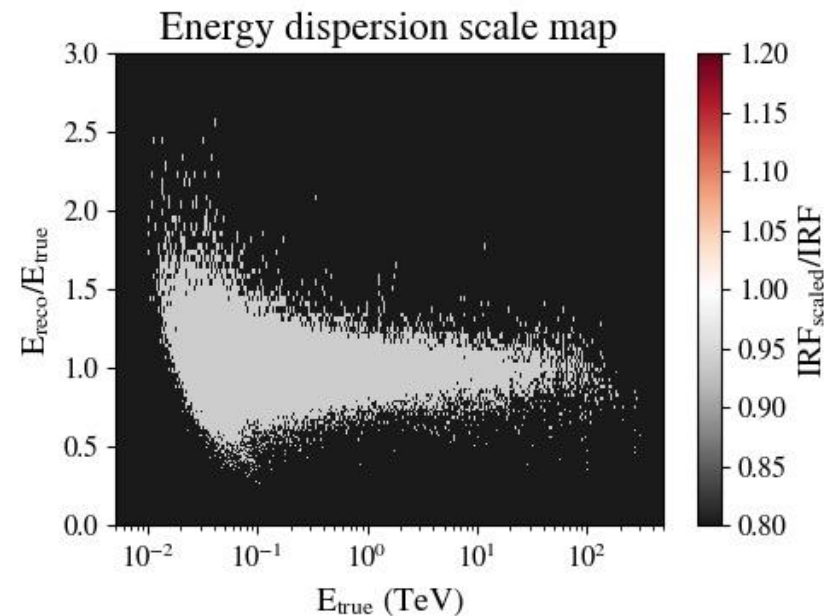
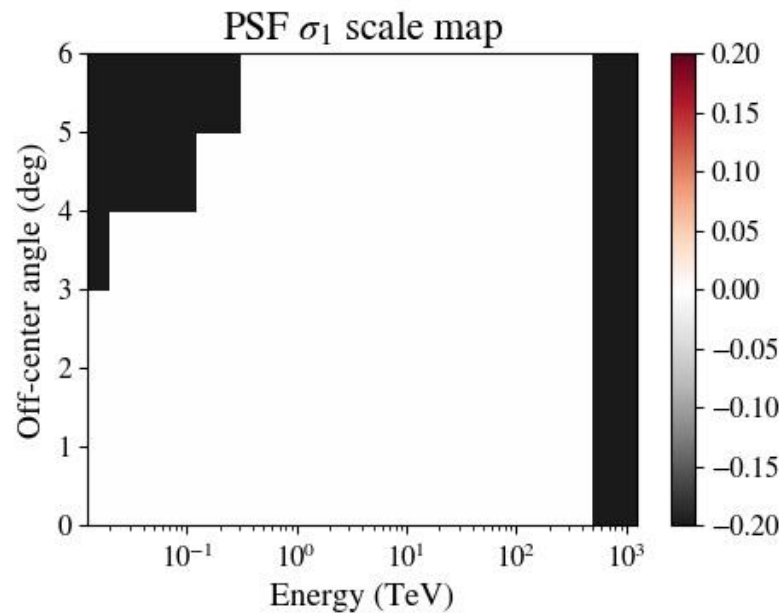
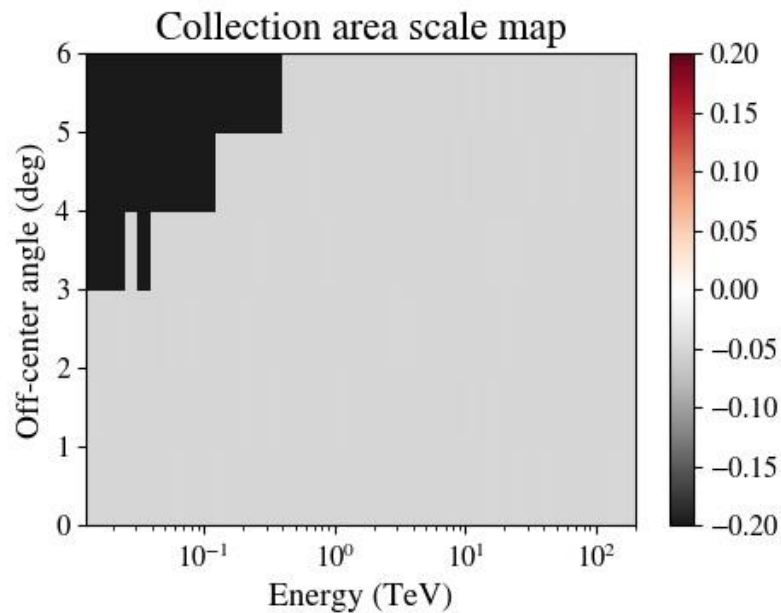


Scaling functions

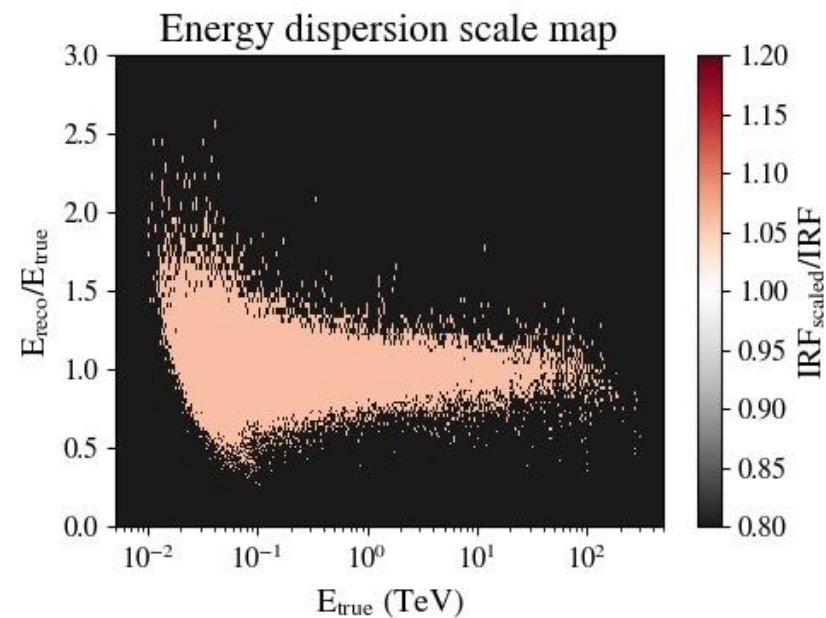
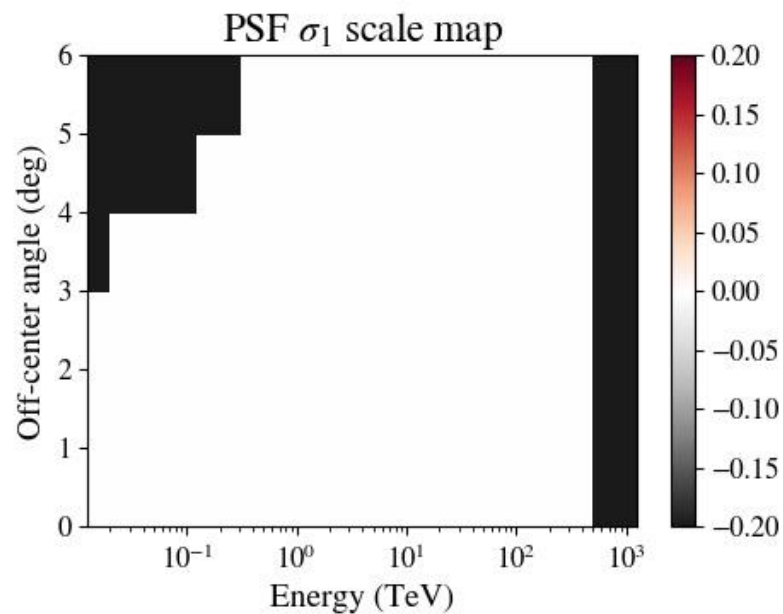
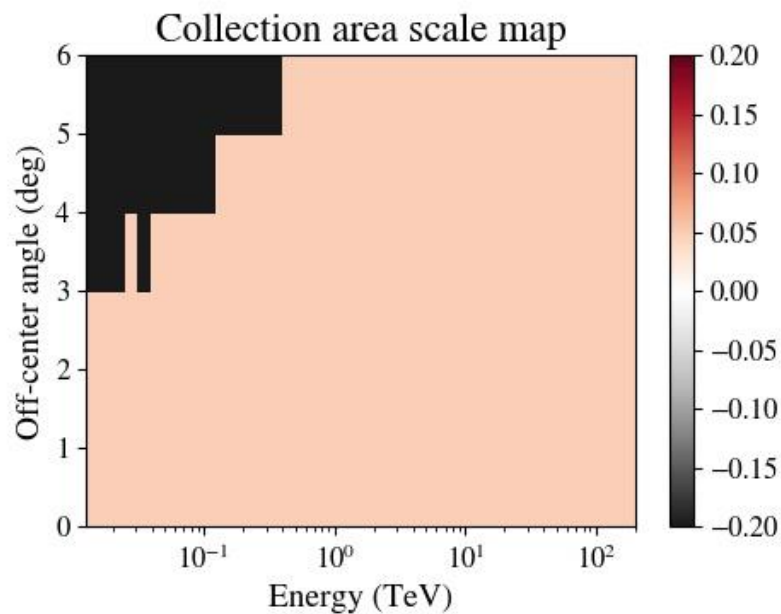
- - and + for each

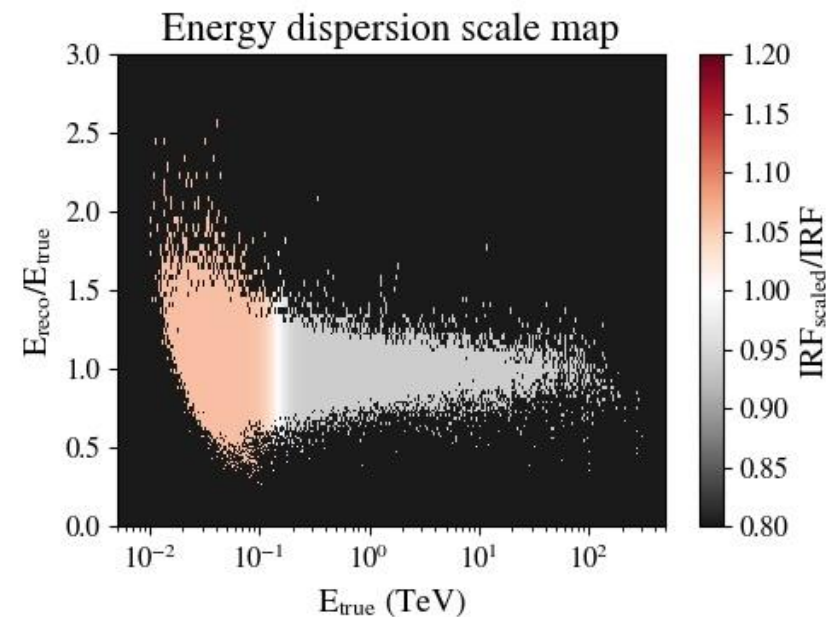
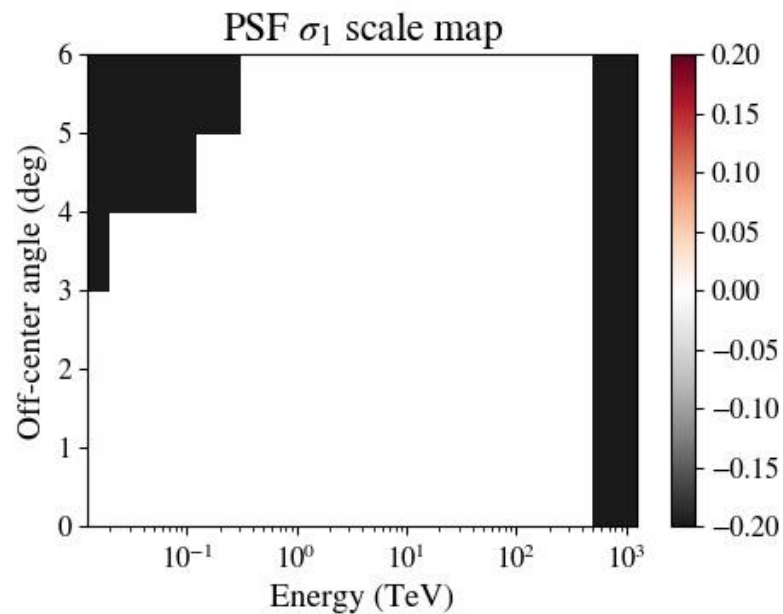
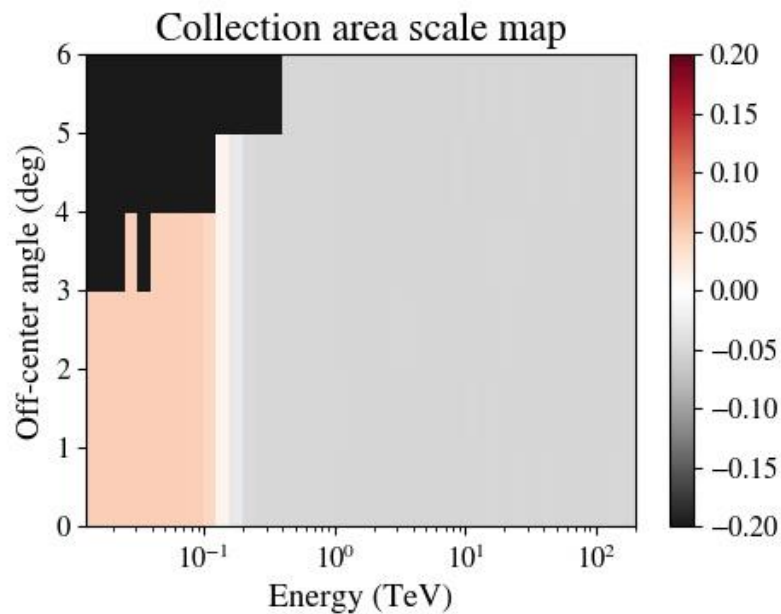
Table 3a: Energy-dependent error functions for CTA North

Modification type	Function, B	Graphics	Applicability
Constant	1		<ul style="list-style-type: none"> • A_{eff}, N: flux normalization • σ_{θ}: small extension • E_{scale}: spectral cut-off • σ_E: search for lines
Gradient	$[\ln(E/E_{\text{min}}) + \ln(E/E_{\text{max}})] / \ln(E_{\text{max}}/E_{\text{min}})$		<ul style="list-style-type: none"> • A_{eff}, N: spectral index, spectral cut-off • E_{scale}: spectral curvature
Step	$\tanh[\ln(E/E_s) / (1.31 \sigma(E_s)/E_s)]$		<ul style="list-style-type: none"> • A_{eff}, N: spectral index, spectral cut-off • E_{scale}: spectral curvature

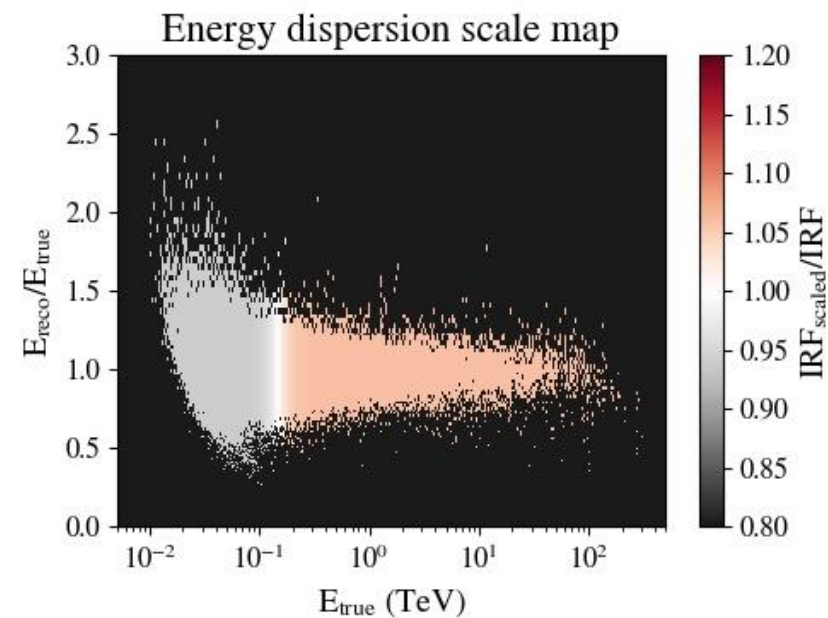
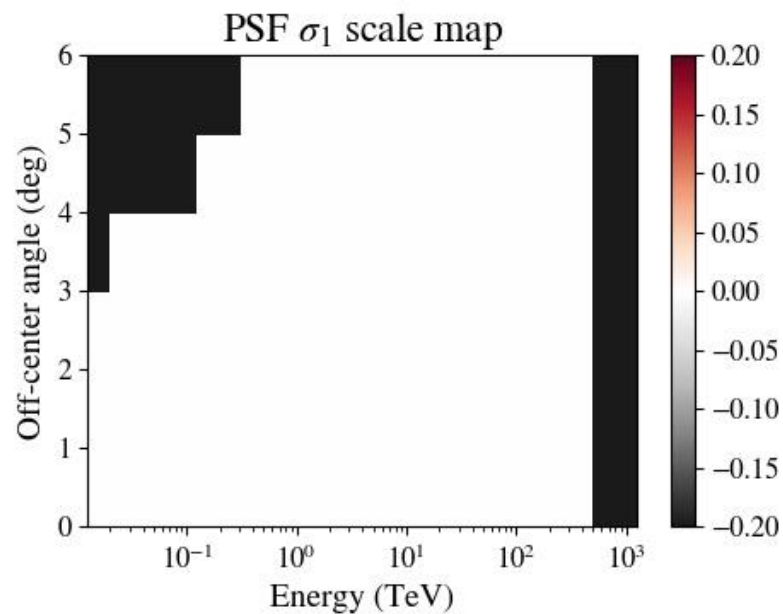
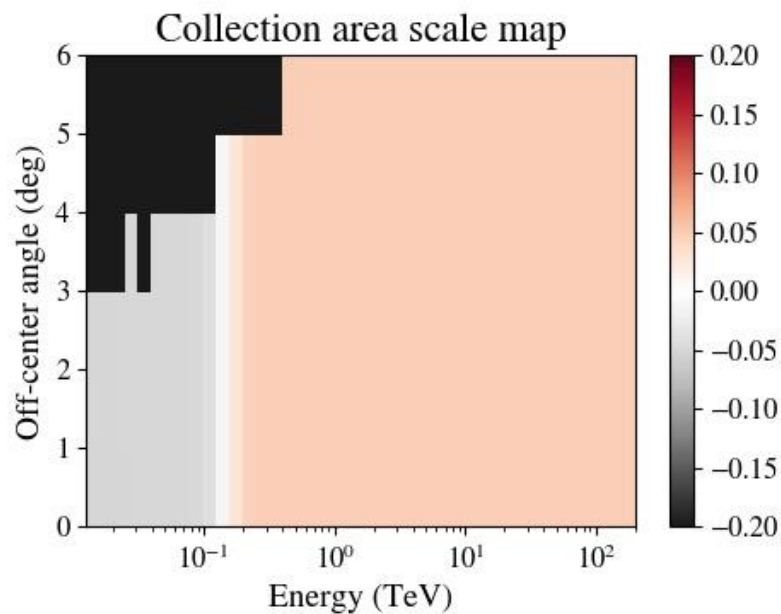


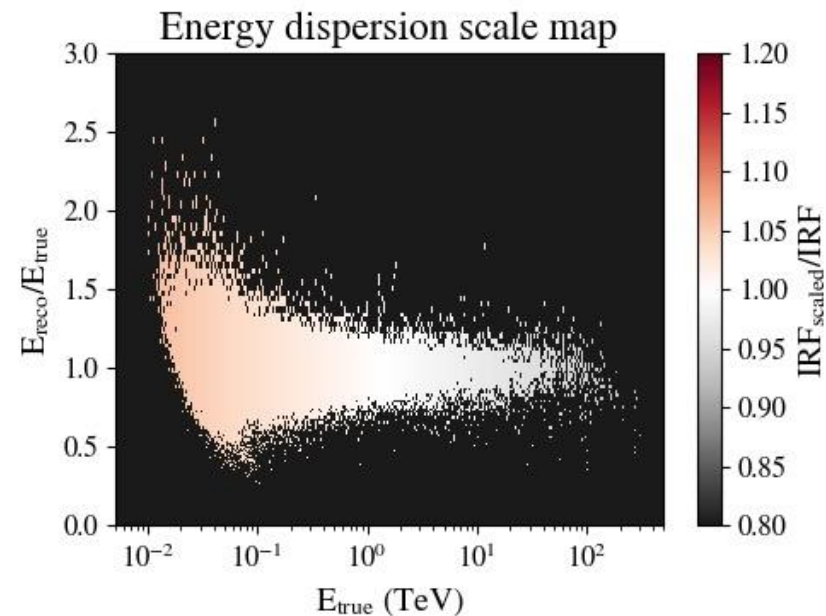
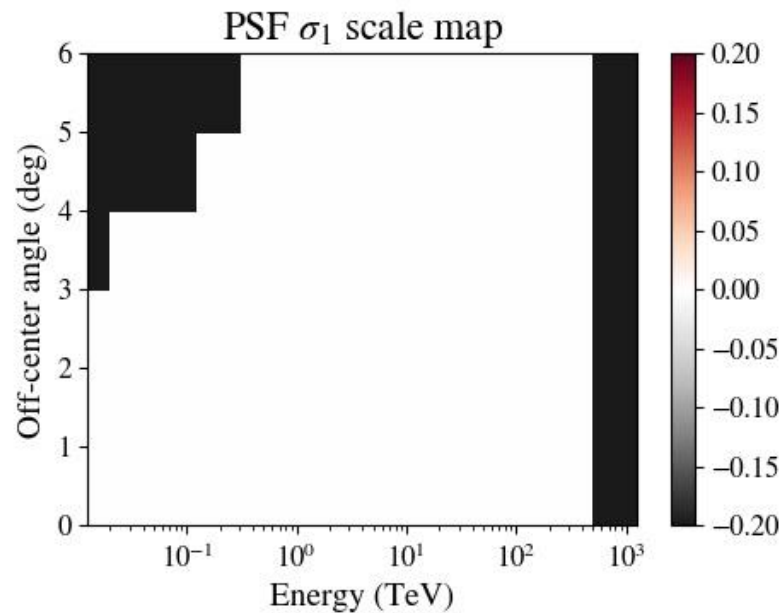
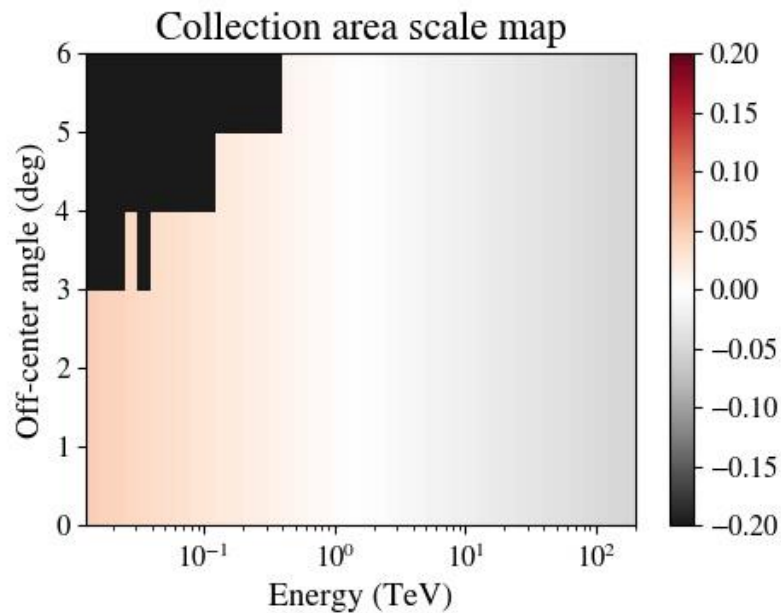
Constant scaling : 5% Aeff, 6% EDisp



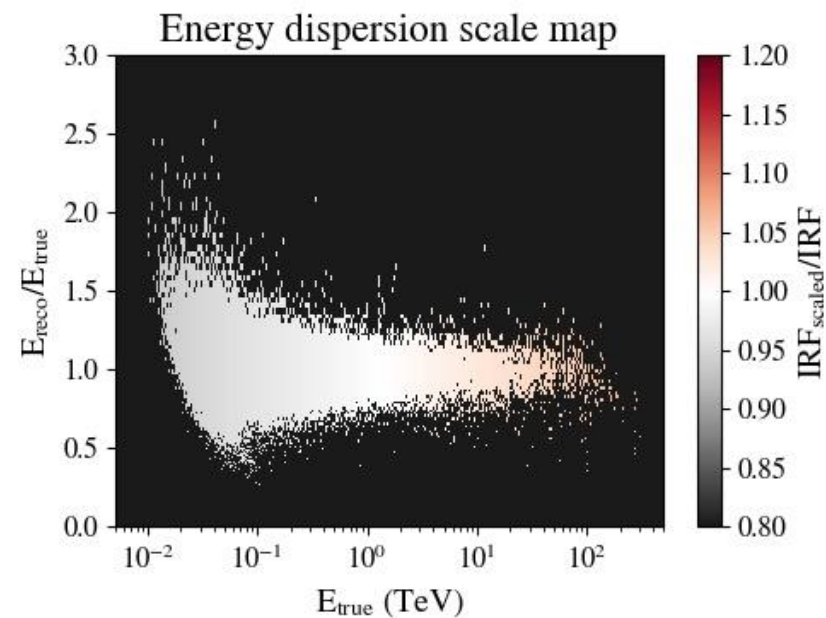
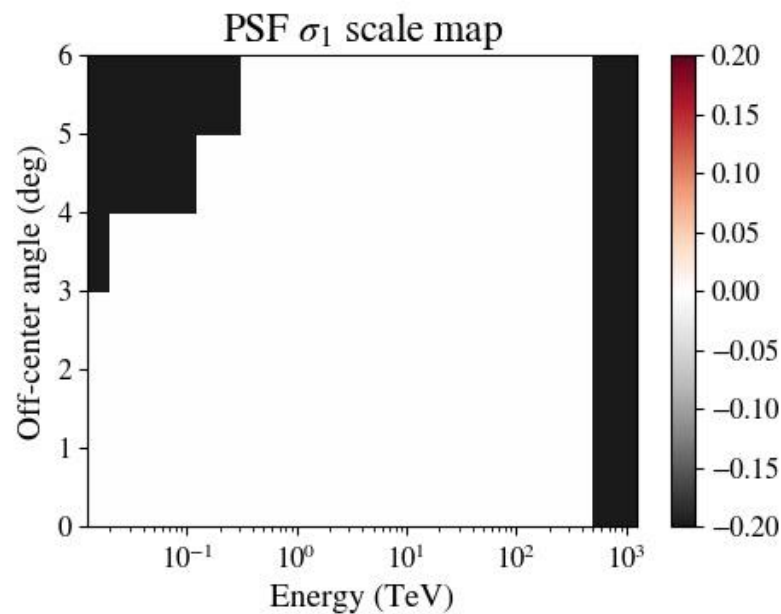
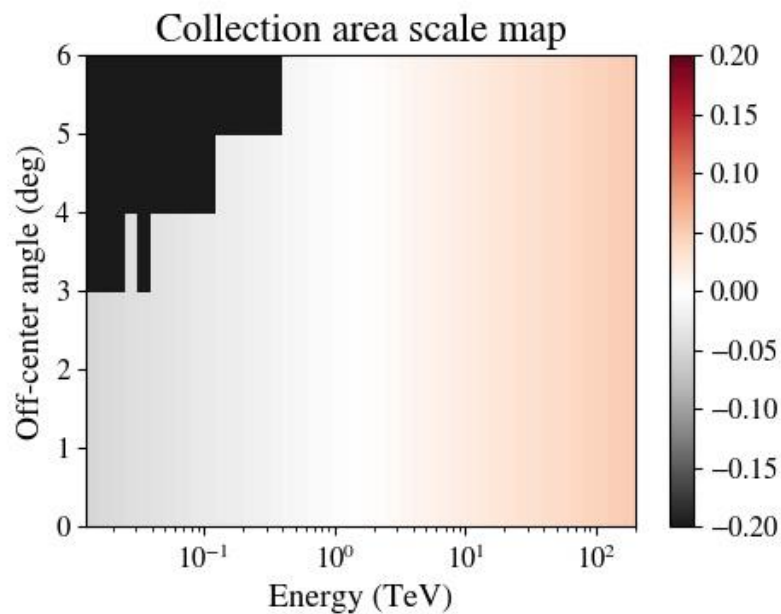


Step scaling : 5% Aeff, 6% EDisp. Break at 0.15TeV (5TeV break will be added later)





Gradient scaling : 5% Aeff, 6% EDisp



Results

<i>Source 1</i> parameters	40mCrab prefactor	-2 index	50TeV cutoff
Observation parameters	40h duration	0.5deg offset	1.6deg FoV
Binning parameters	0.02deg spatial resolution	10 bins per energy decade	

(a) Input parameters: observation data.

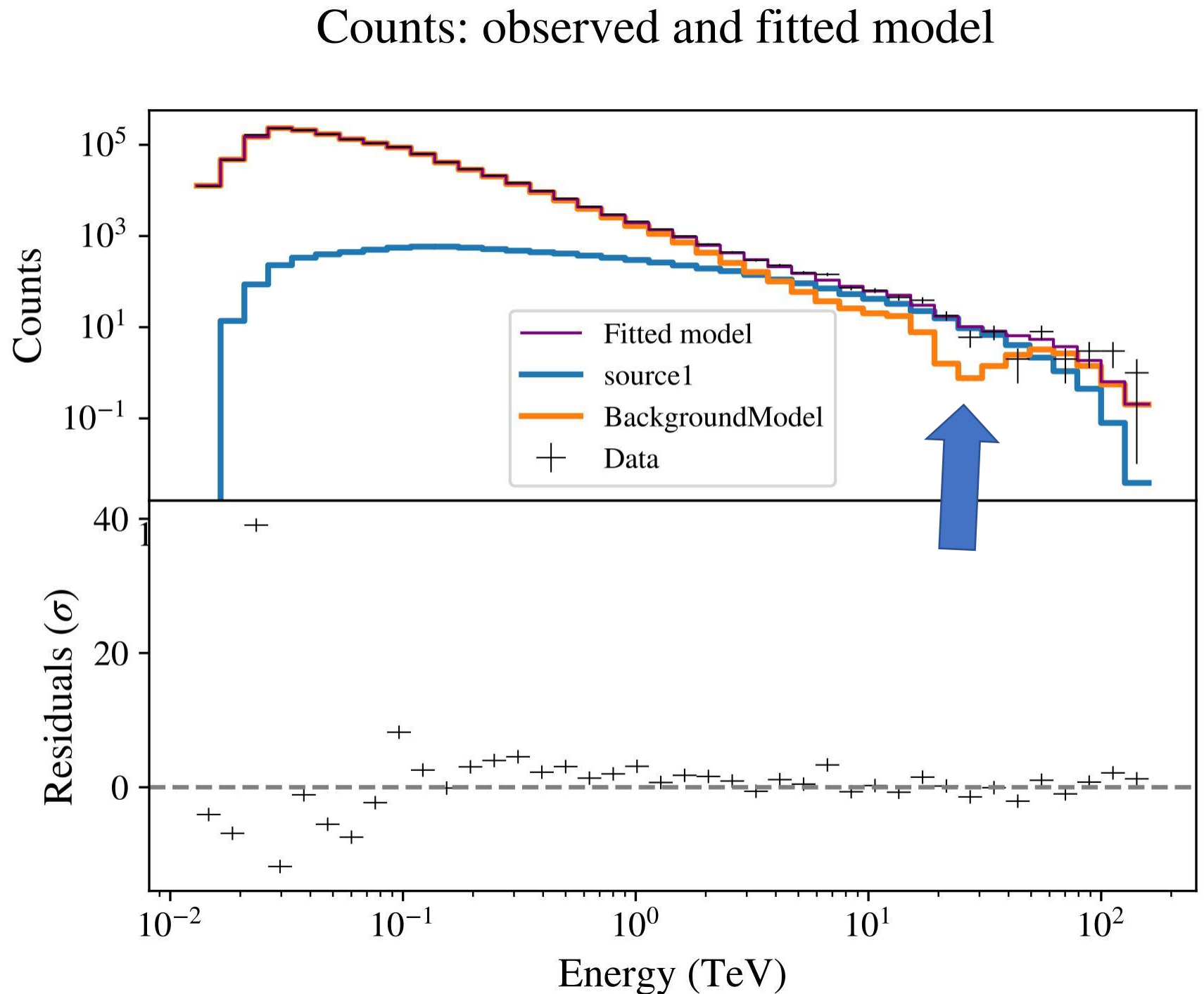
Bracketing	Prefactor (mCrab)	Index	Cutoff (TeV)	Goodness of Fit
None	40.76 ± 0.79	-1.97 ± 0.02	32.89 ± 4.84	
Constant				
⊕	38.53 ± 0.71	-1.99 ± 0.02	73.80 ± 18.95	
⊖	40.39 ± 0.82	-1.97 ± 0.02	48.84 ± 10.93	
Gradient				
⊕	40.09 ± 0.75	-2.00 ± 0.02	52.63 ± 9.87	
⊖	40.90 ± 0.83	-1.97 ± 0.02	29.95 ± 4.53	
Step				
⊕	40.80 ± 0.82	-1.96 ± 0.02	32.53 ± 5.58	
⊖	39.33 ± 0.80	-1.97 ± 0.02	46.4 ± 9.9	

(b) Output parameters: spectral fit. The parameters that do not match the source are highlighted.

Table 1: Parameters for the fitting of observations of *Source 1* with differently bracketed IRFs.

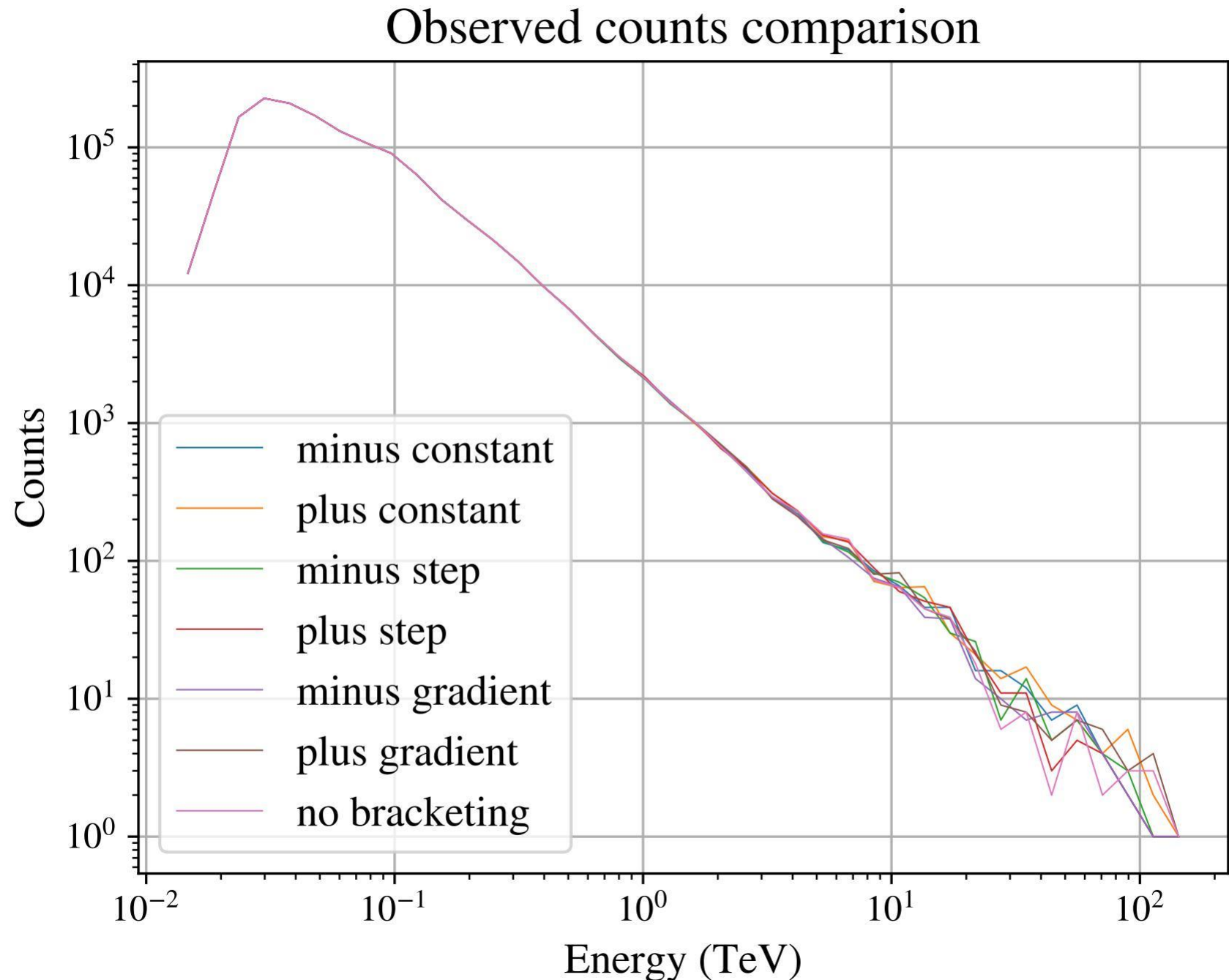
Background problem

- Currently use CTA IRF background
- Should I simulate my own background, close to the IRF's?
- Or keep IRF?
- Or should I switch to On/Off observation?

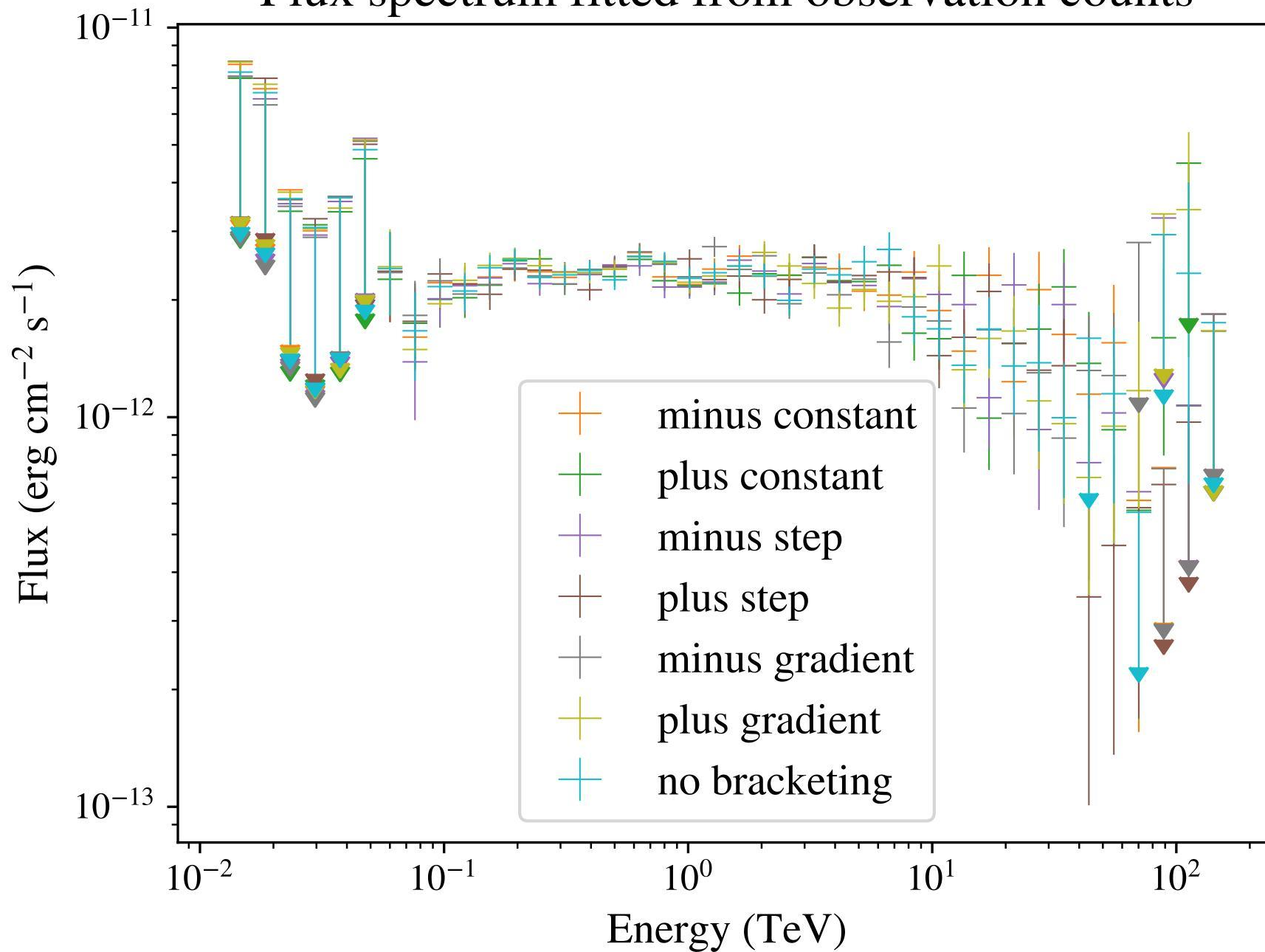


Similarity problem

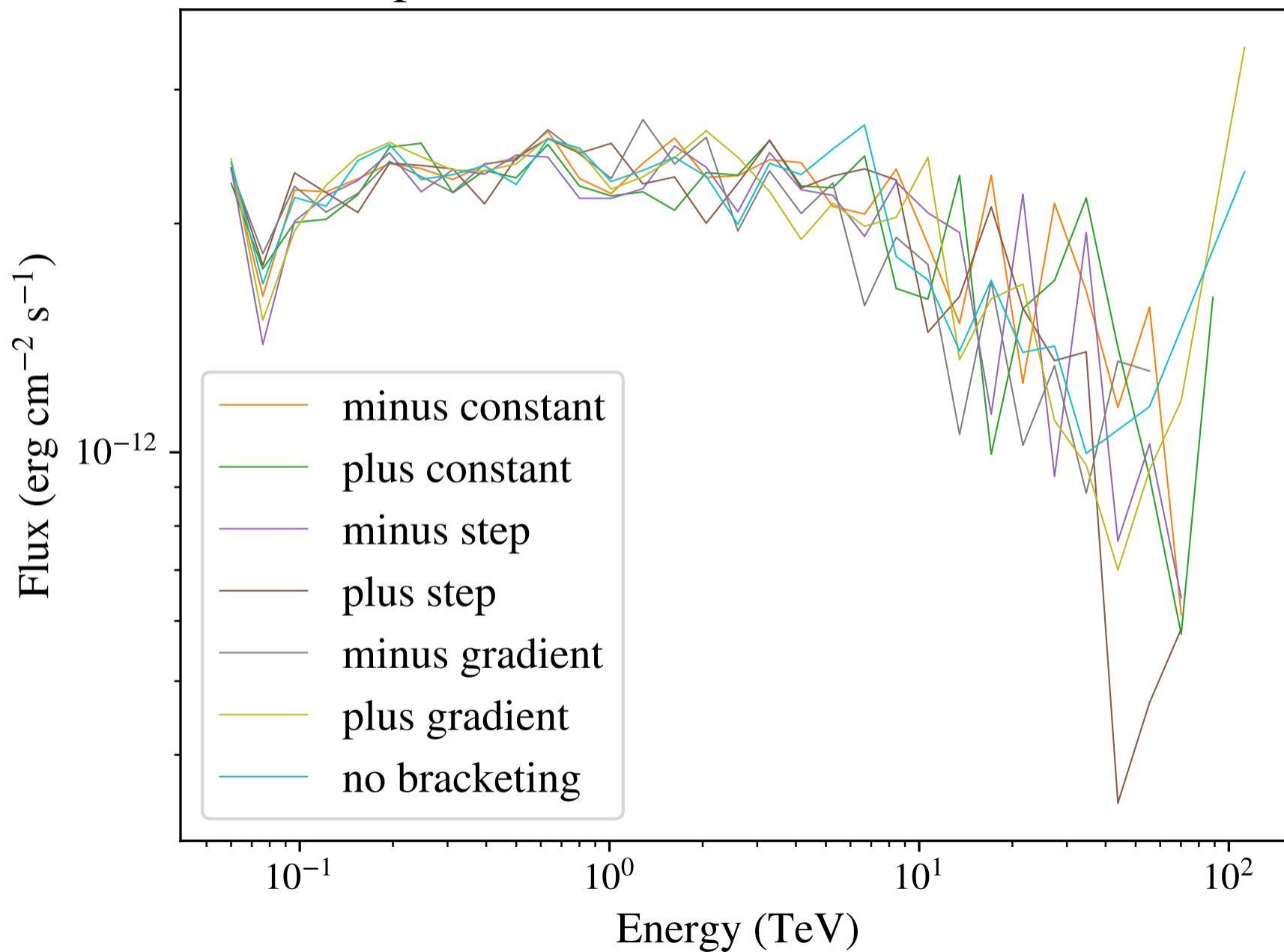
- Why are the simulations/fittings so similar?
- Leads: separating bracketing / exaggerating bracketing / code error ?



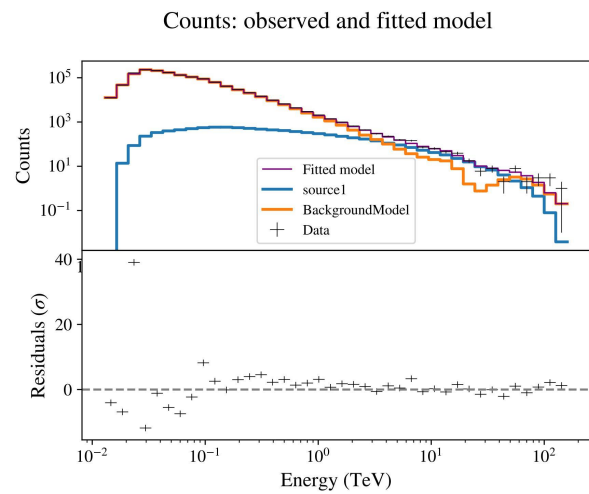
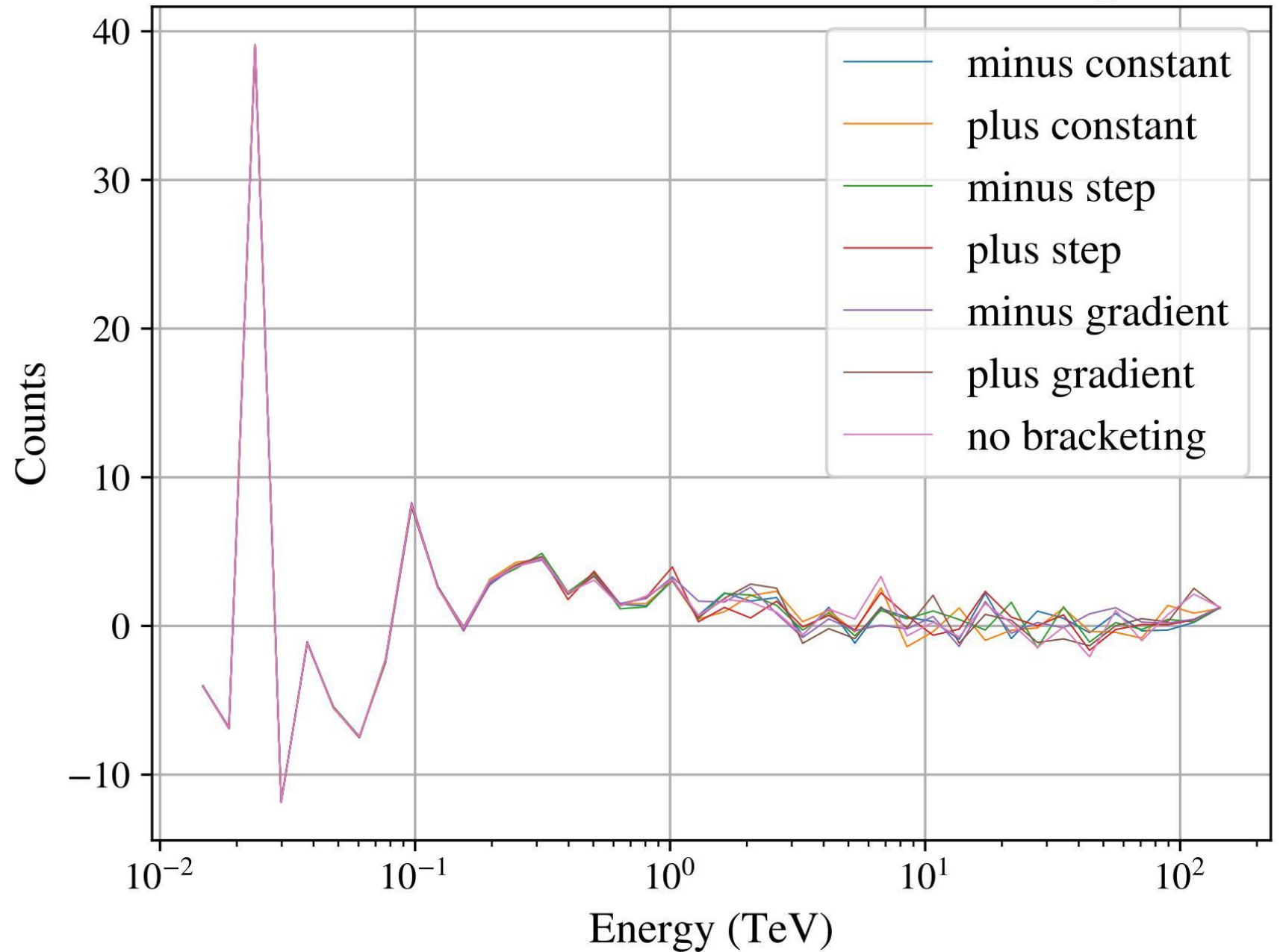
Flux spectrum fitted from observation counts

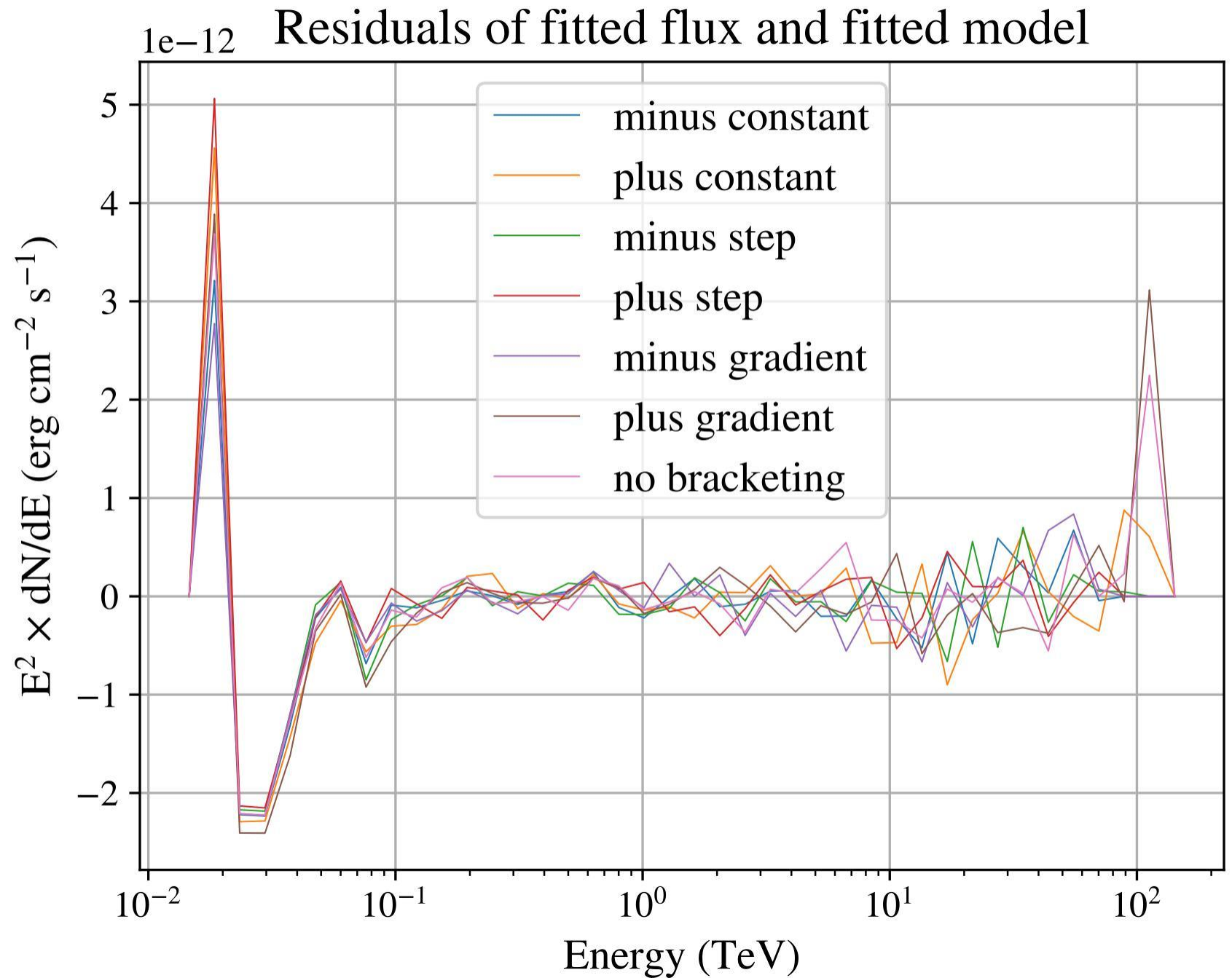
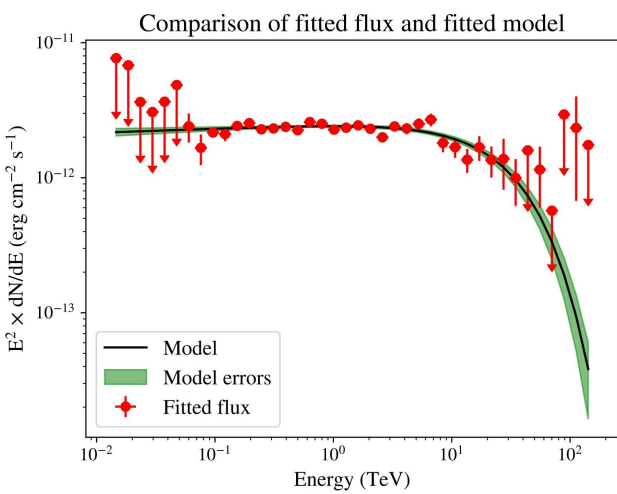


Flux spectrum fitted from observation counts



Residuals of observed and fitted counts comparison





What's next

- New simulations : 20,60,80 mCrab (current 40)
- Separate AEff and EDisp bracketing
- Background change
- Add South breakpoint at 5TeV
- If results+errors aren't too similar, run 100 simulations per set
- Plot flux of each set with instrument sensitivity
- Find out how ctools errors are calculated
- Goodness of fit calculation
- Later: same work with new cutoff
- Possibly: implement resolution bracketing in levgen's function
- Presentation to Working Group 17/07
- End of internship 26/07