Stimulated Neutrino Transitions and Rabi Oscillations

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OUTLINE

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   - What are Neutrinos
   - Neutrino Oscillations
   - Why Do Neutrinos Oscillate
2. Matter Effect
   - Matter Interaction
   - MSW Effect
   - Solar Neutrino Problem
   - Stimulated Neutrino Oscillations
3. Stimulated Neutrino Flavor Transitions
   - Hamiltonian and Basis
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4. Jacobi-Anger Expansion
   - Basis and Formalism
5. Summary

Overview

Introduction
- What are Neutrinos
- Neutrino Oscillations
- Why Do Neutrinos Oscillate

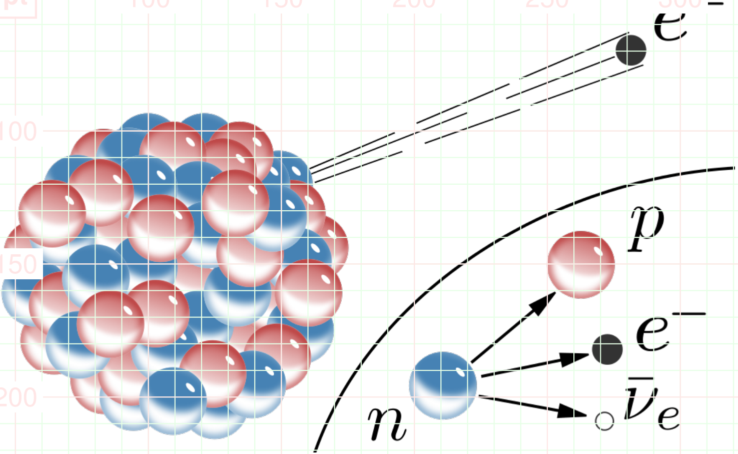
Matter Effect

Stimulated Neutrino Flavor Transitions

Jacobi-Anger Expansion

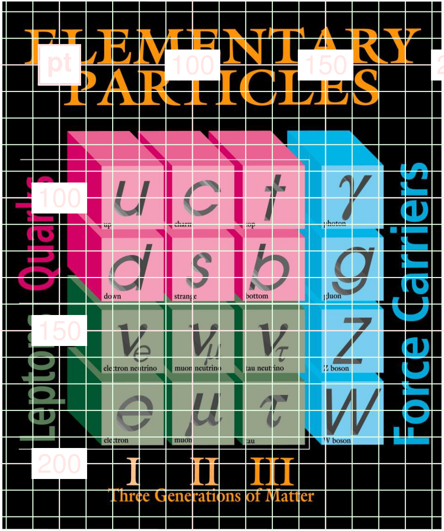
Summary

What are Neutrinos?



Beta decay and antineutrino production. Source: Beta_Decay@Wikipedia

What are Neutrinos?

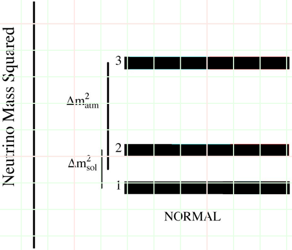


Fermilab 95-759

Table of elementary particles. Source: Fermilab

Neutrinos are

- fermions,
- electrically neutral,
- light.



Adapted from Olga Mena & Stephen Parke (2004)

What are Neutrinos?

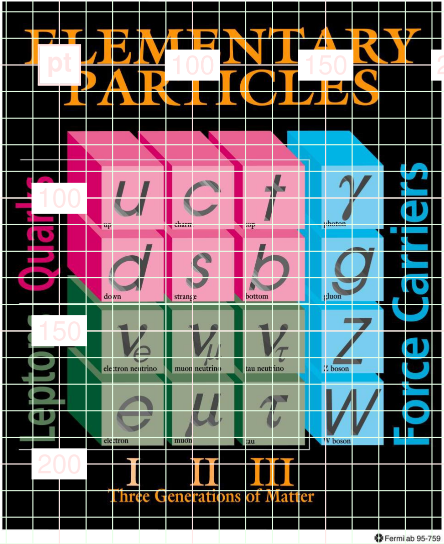


Table of elementary particles. Source: Fermilab

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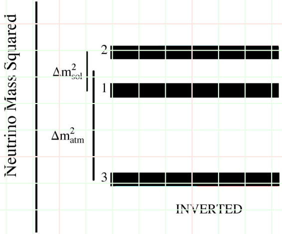

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What is Neutrino Oscillation?

Neutrino Oscillation

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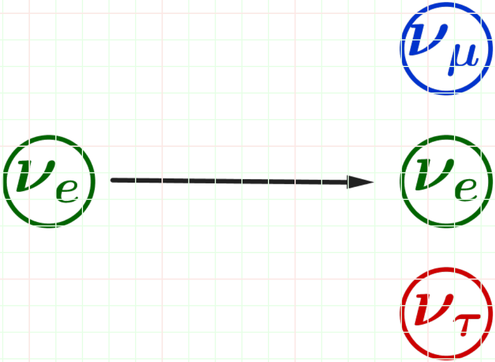
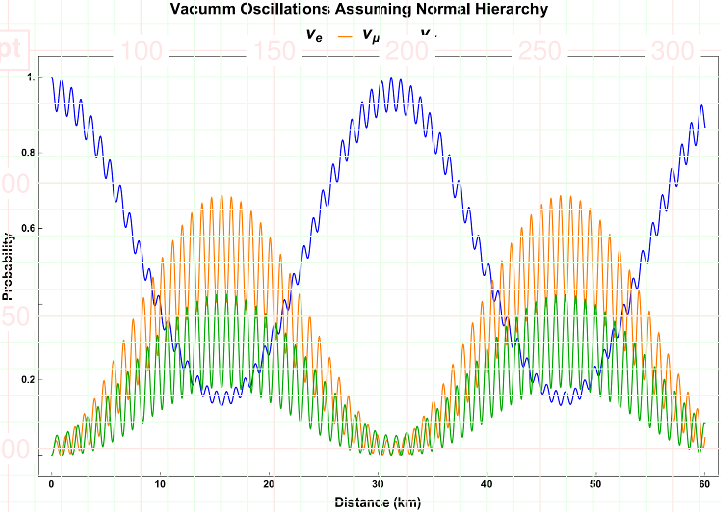
Neutrino Flavor Conversion



Neutrino Oscillations

What is Neutrino Oscillation?



Probabilities of finding neutrinos to be in each flavor.

WHY DO NEUTRINOS OSCILLATE?

Equation of Motion

$$i\partial_x |\Psi\rangle = \hat{\mathbf{H}} |\Psi\rangle$$

- Basis: Hamiltonian diagonalized basis/mass basis/propagation basis, $\{|\nu_1\rangle, |\nu_2\rangle\}$.
- $$\mathbf{H} = -\frac{\omega_v}{2}\sigma_3, \qquad \text{where } \omega_v = \frac{\delta m^2}{2E} = \frac{m_2^2 - m_1^2}{2E}.$$
- The system can be solved given initial condition of the amplitudes of the two eigenstates $(\langle \nu_1|\Psi(0)\rangle, \langle \nu_2|\Psi(0)\rangle)^T$,

$$\begin{pmatrix} \langle \nu_1|\Psi(x)\rangle \\ \langle \nu_2|\Psi(x)\rangle \end{pmatrix} = \begin{pmatrix} \langle \nu_1|\Psi(0)\rangle \exp(i\omega_v x/2) \\ \langle \nu_2|\Psi(0)\rangle \exp(-i\omega_v x/2) \end{pmatrix}$$

Why Do Neutrinos Oscillate?

Flavor basis

Neutrino wave function in flavor basis $\{|\nu_e\rangle, |\nu_\mu\rangle\}$ is related to state in energy basis $\{|\nu_1\rangle, |\nu_2\rangle\}$ through

$$\begin{pmatrix} |\nu_e\rangle \\ |\nu_\mu\rangle \end{pmatrix} = \begin{pmatrix} \cos\theta_v & \sin\theta_v \\ -\sin\theta_v & \cos\theta_v \end{pmatrix} \begin{pmatrix} |\nu_1\rangle \\ |\nu_2\rangle \end{pmatrix}$$

θ_v: vacuum mixing angle

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θ_v: vacuum mixing angle

Hamiltonian H

Mass basis

$$\frac{\omega_v}{2}\begin{pmatrix} -1 & 0 \\ 0 & 1 \end{pmatrix} = -\frac{\omega_v}{2}\boldsymbol{\sigma_3}$$

Flavor basis

$$\frac{\omega_v}{2}\begin{pmatrix} -\cos 2\theta_v & \sin 2\theta_v \\ \sin 2\theta_v & \cos 2\theta_v \end{pmatrix} = \frac{\omega_v}{2}\left(-\cos 2\theta_v \boldsymbol{\sigma_3} + \sin 2\theta_v \boldsymbol{\sigma_1}\right)$$

Nature of Neutrino Oscillation

pt 100 150 200 250 300 350

Transition Probability

100

$$P(|\nu_e\rangle \to |\nu_\mu\rangle) = \sin^2(2\theta_v)\sin^2(\omega_v x/2)$$

150

- $\omega_v = (m_2^2 - m_1^2)/2E$ determines oscillation wavelength.
- Mixing angle θ_v determines flavor oscillation amplitude.

200

250

Overview

Introduction

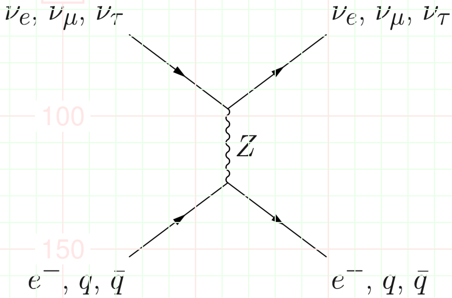
Matter Effect
- Matter Interaction
- MSW Effect
- Solar Neutrino Problem
- Stimulated Neutrino Oscillations

Stimulated Neutrino Flavor Transitions

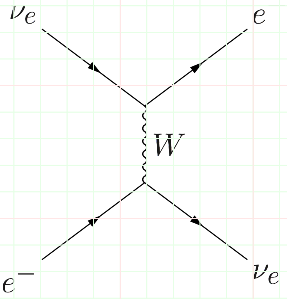
Jacobi-Anger Expansion

Summary

Matter Interaction



Neutral current interaction between $\nu_\mathrm{e}, \nu_\mu, \nu_\tau$,
and quarks and antiquarks.



Charged current interaction between ν_e and e^-

Matter Interaction

Hamiltonian with matter interaction in flavor basis ($\omega_{\rm v} = \delta m^2/2E$):

$$\mathbf{H} = \frac{\omega_{\rm v}}{2}\begin{pmatrix} -\cos 2\theta_{\rm v} & \sin 2\theta_{\rm v} \\ \sin 2\theta_{\rm v} & \cos 2\theta_{\rm v} \end{pmatrix} \pm \sqrt{2} G_{\rm F} n_{\rm e}(x) \begin{pmatrix} 1 & 0 \\ 0 & 0 \end{pmatrix}$$

- Vacuum Hamiltonian
- Matter interaction

MATTER INTERACTION

Hamiltonian with matter interaction in flavor basis ($\omega_{\mathrm{v}} = \delta m^2/2E$):

$$\mathbf{H} = \frac{\omega_{\mathrm{v}}}{2}\left(-\cos 2\theta_{\mathrm{v}}\boldsymbol{\sigma_3} + \sin 2\theta_{\mathrm{v}}\boldsymbol{\sigma_1}\right) + \frac{\lambda(x)}{2}\boldsymbol{\sigma_3}$$

- Vacuum Hamiltonian
- Matter interaction
- $\lambda(x) = \sqrt{2}G_{\mathrm{F}}n_{\mathrm{e}}(x)$

MSW Effect

Hamiltonian in Vacuum

$$\mathbf{H}_{\text{vacuum}} = -\frac{\omega_{\text{v}} \cos 2\theta_{\text{v}}}{2}\boldsymbol{\sigma}_\mathbf{3} + \frac{\omega_{\text{v}} \sin 2\theta_{\text{v}}}{2}\boldsymbol{\sigma}_\mathbf{1}$$

$$\begin{aligned}\mathbf{H} &= \frac{\lambda(x) - \omega_{\text{v}} \cos 2\theta_{\text{v}}}{2}\boldsymbol{\sigma}_\mathbf{3} + \frac{\omega_{\text{v}} \sin 2\theta_{\text{v}}}{2}\boldsymbol{\sigma}_\mathbf{1} \\ &= \frac{\omega_{\text{m}}(x) \cos 2\theta_{\text{m}}(x)}{2}\boldsymbol{\sigma}_\mathbf{3} + \frac{\omega_{\text{m}}(x) \sin 2\theta_{\text{m}}(x)}{2}\boldsymbol{\sigma}_\mathbf{1},\end{aligned}$$

where

$$\begin{aligned}\omega_{\text{m}}(x) &= \sqrt{(\lambda(x) - \omega_{\text{v}} \cos 2\theta_{\text{v}})^2 + \omega_{\text{v}}^2 \sin^2 2\theta_{\text{v}}}, \\ \tan 2\theta_{\text{m}}(x) &= \frac{\omega_{\text{v}} \sin 2\theta_{\text{v}}}{\omega_{\text{v}} \cos 2\theta_{\text{v}} - \lambda(x)}.\end{aligned}$$

MSW Effect

Constant matter profile λ_0 as an example,

Significance of θ_m

Define matter basis $\{|\nu_\mathrm{L}\rangle, |\nu_\mathrm{H}\rangle\}$

$$\begin{pmatrix} |\nu_e\rangle \\ |\nu_\mu\rangle \end{pmatrix} = \begin{pmatrix} \cos\theta_\mathrm{m} & \sin\theta_\mathrm{m} \\ -\sin\theta_\mathrm{m} & \cos\theta_\mathrm{m} \end{pmatrix} \begin{pmatrix} |\nu_\mathrm{L}\rangle \\ |\nu_\mathrm{H}\rangle \end{pmatrix}$$

In matter basis

$$\mathbf{H}_{\text{matter-basis}} = -\frac{\omega_\mathrm{m}}{2}\boldsymbol{\sigma_3}$$

MSW Resonance

Hamiltonian with Matter Potential

$$\mathbf{H} = \frac{\lambda(x) - \omega_\text{v}\cos 2\theta_\text{v}}{2}\boldsymbol{\sigma_3} + \frac{\omega_\text{v}\sin 2\theta_\text{v}}{2}\boldsymbol{\sigma_1}$$
$$= \frac{\omega_\text{m}(x)\cos 2\theta_\text{m}(x)}{2}\boldsymbol{\sigma_3} + \frac{\omega_\text{m}(x)\sin 2\theta_\text{m}(x)}{2}\boldsymbol{\sigma_1}$$

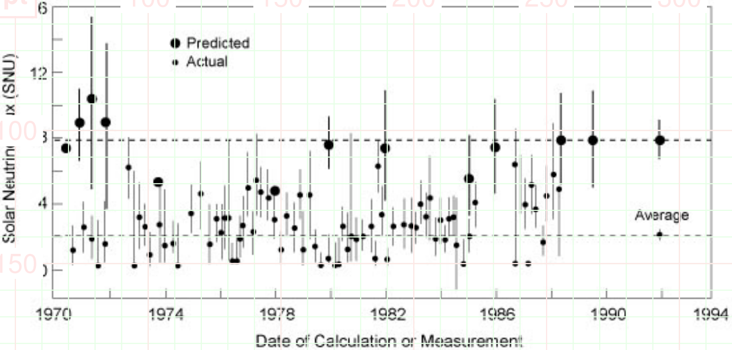
$$\tan 2\theta_\text{m}(x) = \frac{\omega_\text{v}\sin 2\theta_\text{v}}{\omega_\text{v}\cos 2\theta_\text{v} - \lambda(x)}.$$

$$\begin{pmatrix} |\nu_e\rangle \\ |\nu_\mu\rangle \end{pmatrix} = \begin{pmatrix} \cos\theta_\text{m} & \sin\theta_\text{m} \\ -\sin\theta_\text{m} & \cos\theta_\text{m} \end{pmatrix} \begin{pmatrix} |\nu_\text{L}\rangle \\ |\nu_\text{H}\rangle \end{pmatrix}$$

Transition Probability

$$P(|\nu_e\rangle \to |\nu_\mu\rangle) = \sin^2(2\theta_\text{m})\sin^2(\omega_\text{m} x)$$

Solar Neutrino Problem



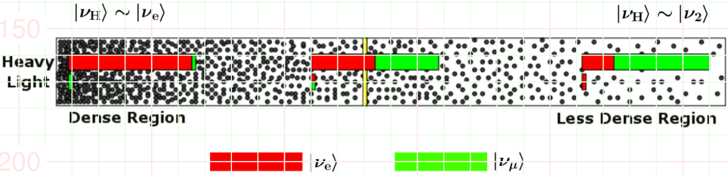
Chlorine detector (Homestake experiment) results and theory predictions. SNU: 1 event for 10^{36} target atoms per second. Kenneth R. Lang (2010)

MSW Effect and Solar Neutrinos

$$= \frac{\lambda(x) - \omega_{\rm v} \cos 2\theta_{\rm v}}{2} - \frac{\omega_{\rm v} \sin 2\theta_{\rm v}}{2} \sigma_1$$

$$\begin{pmatrix} |\nu_{\rm L}\rangle \\ |\nu_{\rm H}\rangle \end{pmatrix} = \begin{pmatrix} \cos\theta_{\rm m} & -\sin\theta_{\rm m} \\ \sin\theta_{\rm m} & \cos\theta_{\rm m} \end{pmatrix} \begin{pmatrix} |\nu_{\rm e}\rangle \\ |\nu_{\mu}\rangle \end{pmatrix}$$

$$\mathbf{H}_{\text{matter-basis}} = -\frac{\omega_{\rm m}}{2} \boldsymbol{\sigma_3}$$



Yellow bar is the resonance point. Red: $|\nu_e\rangle$. Green: $|\nu_\mu\rangle$. Adapted from Smirnov, 2003.

MSW Effect

Suppose $\omega_\mathrm{v} = (m_2^2 - m_1^2)/2E < 0$,

$$\mathbf{H} = -\frac{\omega_\mathrm{v}}{2}\begin{pmatrix} -\cos 2\theta_\mathrm{v} & \sin 2\theta_\mathrm{v} \\ \sin 2\theta_\mathrm{v} & \cos 2\theta_\mathrm{v} \end{pmatrix} + \sqrt{2}G_\mathrm{F} n_\mathrm{e}(x)\begin{pmatrix} 1 & 0 \\ 0 & 0 \end{pmatrix}$$

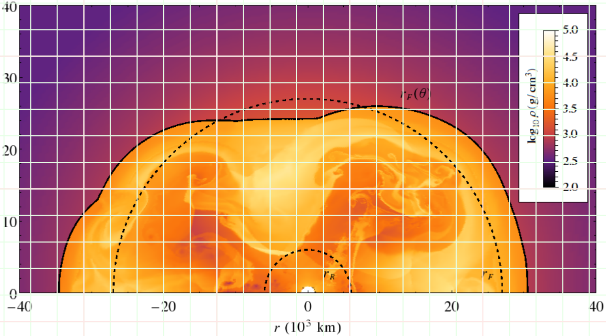
$$\downarrow$$

$$\mathbf{H} = \left(\frac{-\omega_\mathrm{v}}{2}\cos 2\theta_\mathrm{v} + \frac{\lambda(x)}{2}\right)\boldsymbol{\sigma}_3 - \frac{\omega_\mathrm{v}}{2}\sin 2\theta_\mathrm{v}\boldsymbol{\sigma}_1$$

Supernova Matter Density Profile

Why Do We Care

Astrophysical environments: supernovae, accretion disks etc



Supernova shock and turbulence. E. Borriello, et al (2014)

$$\Delta n_e(r) = \sum_n c_n \sin(k_n r + \phi_n)$$

Matter Profile

$$\lambda(x) = \lambda_0 + \delta\lambda(x)$$

Basis

Background matter basis: Hamiltonian is diagonalized with only background matter profile λ_0,

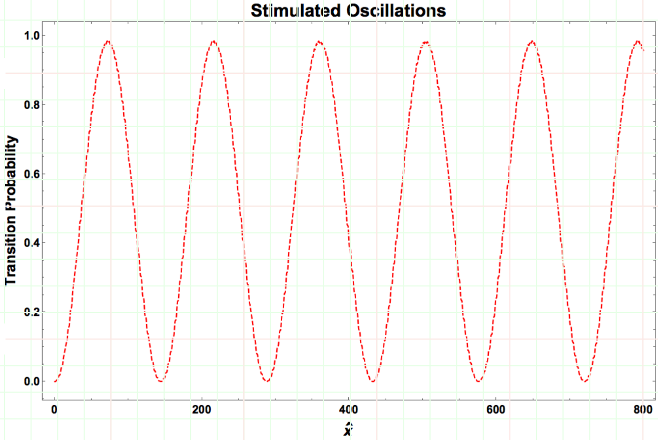
$$\mathrm{H}_{\mathrm{background}} = -\frac{\omega_{\mathrm{m}}}{2}\boldsymbol{\sigma_3}.$$

Hamiltonian

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \delta\lambda(x)\cos 2\theta_{\mathrm{m}}\right)\boldsymbol{\sigma_3} - \frac{\delta\lambda(x)}{2}\sin\theta_{\mathrm{m}}\boldsymbol{\sigma_1}.$$

STIMULATED NEUTRINO OSCILLATIONS

P. Krastev and A. Smirnov (1989); J. Kneller et al (2013); Patton et al (2014);



Stimulated oscillations. $\lambda(x) = \lambda_0 + A\sin(kx)$ with $\hat{x} = \omega_{\rm m} x$, $A = 0.1\omega_{\rm m}$, $k = 0.995\omega_{\rm m}$, $\theta_{\rm m} = \pi/6$

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Hamiltonian

pt 100 150 200 250 300 350

Hamiltonian in Background Matter Basis

100

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \delta\lambda(x)\cos 2\theta_{\mathrm{m}}\right)\sigma_3 - \frac{\delta\lambda(x)}{2}\sin 2\theta_{\mathrm{m}}\sigma_1 .$$

Matter profile

150

$$\lambda(x) = \lambda_0 + A\cos(kx),$$

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \cos 2\theta_{\mathrm{m}} A\cos(kx)\right)\sigma_3 - \frac{\sin 2\theta_{\mathrm{m}}}{2} A\cos(kx)\sigma_1 .$$

200

250

Rabi Oscillations

Rabi Oscillation

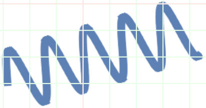
Hamiltonian

$$-\frac{\omega}{2}\sigma_3 - \frac{\alpha}{2}\begin{pmatrix} 0 & e^{ikt} \\ e^{-ikt} & 0 \end{pmatrix}$$

$E_2 = \frac{\omega_0}{2}$

$E_1 = -\frac{\omega_0}{2}$

Incoming light



Frequency : k

Rabi Oscillations

$$\frac{1}{2}\left(-\omega_{\mathrm{m}} + \cos 2\theta_{\mathrm{m}} A\cos(kx)\right)\sigma_3 - \frac{\sin 2\theta_{\mathrm{m}}}{2} A\cos(kx)\sigma_1$$

R Oscillat

Hamiltonian

$$-\frac{\omega_0}{2}\sigma_3 - \frac{\alpha}{2}\begin{pmatrix} 0 & e^{ikt} \\ e^{-ikt} & 0 \end{pmatrix}$$

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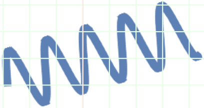


Frequency : k

The transition probability from low energy to high energy is

$$P_{1\to 2} = \frac{\alpha^2}{\alpha^2 + (\omega_0 - k)^2}\sin^2\left(\frac{\Omega_{\mathrm{R}}}{2}t\right),$$

wh

$$\Omega_{\mathrm{R}} = \sqrt{\alpha^2 + (\omega_0 - k)^2}$$

is Rabi frequency.

Rabi Oscillations

Rabi Oscillations

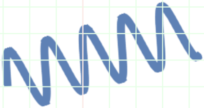
Hamiltonian

$$-\frac{\omega_0}{2}\sigma_3 - \frac{\alpha}{2}\begin{pmatrix} 0 & e^{ikt} \\ e^{-ikt} & 0 \end{pmatrix}$$

$E_2 = \frac{\omega_0}{2}$

$E_1 = -\frac{\omega_0}{2}$

Incoming light



Frequency : k

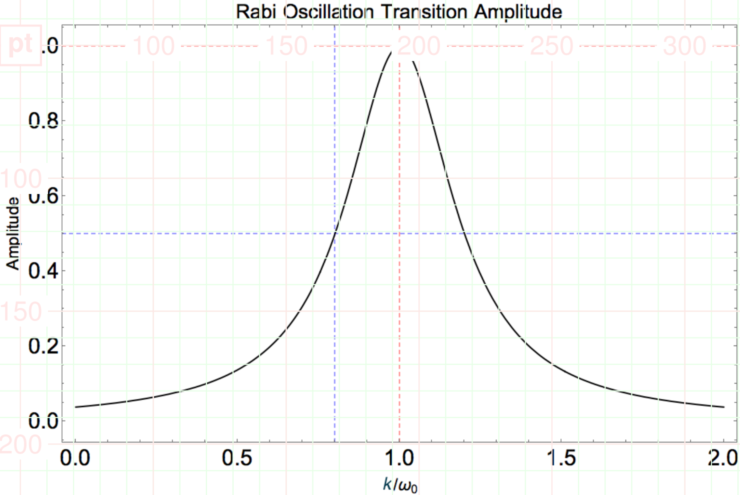
The transition probability from low energy to high energy is

$$P_{1\to 2} = \frac{1}{1+D^2}\sin^2\left(\frac{\Omega_R}{2}t\right),$$

where

$$D = \left|\frac{\omega_0 - k}{\alpha}\right|.$$

Rabi Oscillations



Amplitude of Rabi oscillations for different driving field frequency k

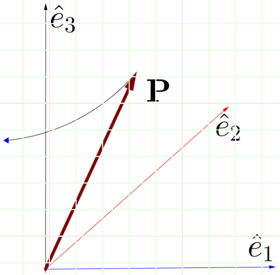
Visualizing Rabi Oscillations

$$-\frac{\omega_0}{2}\sigma_3 - \frac{\alpha}{2}\begin{pmatrix} 0 & e^{ikt} \\ e^{-ikt} & 0 \end{pmatrix}$$

$$-\frac{\omega_0}{2}\sigma_3 - \frac{\alpha}{2}\cos(kt)\sigma_1 + \frac{\alpha}{2}\sin(kt)\sigma_2$$

$$= \begin{pmatrix} \alpha\cos(kt) & -\alpha\sin(kt) & \omega_0 \end{pmatrix} \begin{pmatrix} -\sigma_1/2 \\ -\sigma_2/2 \\ -\sigma_3/2 \end{pmatrix}$$

$$= \vec{H} \cdot (-\vec{\sigma}/2)$$



$$D = \left| \frac{\omega_0 - k}{\alpha} \right|$$

is ratio of the energy gap in corotating frame to width of resonance.

INTERFERENCES OF RABI OSCILLATIONS

$$\begin{pmatrix} 0 \\ 0 \\ \omega_0 \end{pmatrix} + \alpha \begin{pmatrix} \cos(kt) \\ -\sin(kt) \\ 0 \end{pmatrix}$$

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \cos 2\theta_{\mathrm{m}} A\cos(kx)\right)\sigma_3 - \frac{\sin 2\theta_{\mathrm{m}}}{2} A\cos(kx)\sigma_1$$

$$\rightarrow -\frac{\omega_m}{2}\sigma_3 - \frac{A\sin 2\theta_{\mathrm{m}}}{2}\cos(kx)\sigma_1$$

$$\rightarrow \begin{pmatrix} 0 \\ 0 \\ \omega_m \end{pmatrix} + \frac{A\sin 2\theta_{\mathrm{m}}}{2}\begin{pmatrix} \cos(kx) \\ -\sin(kx) \\ 0 \end{pmatrix} + \frac{A\sin 2\theta_{\mathrm{m}}}{2}\begin{pmatrix} \cos(-kx) \\ -\sin(-kx) \\ 0 \end{pmatrix}$$

Two frequencies!

Interferences of Rabi Oscillations

$$\vec{H} = \begin{pmatrix} 0 \\ 0 \\ \omega_m \end{pmatrix} + \alpha_1 \begin{pmatrix} \cos(k_1 x) \\ -\sin(k_1 x) \\ 0 \end{pmatrix} + \alpha_2 \begin{pmatrix} \cos(k_2 x) \\ -\sin(k_2 x) \\ 0 \end{pmatrix}$$

Corotating frame of the second frequency,

$$\vec{H} = \begin{pmatrix} 0 \\ 0 \\ \omega_m - k_2 \end{pmatrix} + \alpha_1 \begin{pmatrix} \cos(k_1 - k_2 x) \\ -\sin(k_1 - K_2 x) \\ 0 \end{pmatrix} + \alpha_2 \begin{pmatrix} 1 \\ 0 \\ 0 \end{pmatrix}$$

Energy gap in this frame becomes the length of the vector

$$\begin{pmatrix} 0 \\ 0 \\ \omega_m - k_2 \end{pmatrix} + \alpha_2 \begin{pmatrix} 1 \\ 0 \\ 0 \end{pmatrix}$$

INTERFERENCES OF RABI OSCILLATIONS

$$\vec{H} = \begin{pmatrix} 0 \\ 0 \\ \omega_m \end{pmatrix} + \alpha_1 \begin{pmatrix} \cos(k_1 x) \\ \sin(k_1 x) \\ 0 \end{pmatrix} + \alpha_2 \begin{pmatrix} \cos(k_2 x) \\ \sin(k_2 x) \\ 0 \end{pmatrix}$$

Corotating frame of the second frequency,

$$\vec{H} = \begin{pmatrix} 0 \\ 0 \\ \omega_m - k_2 \end{pmatrix} + \alpha_1 \begin{pmatrix} \cos(k_1 - k_2 x) \\ -\sin(k_1 - K_2 x) \\ 0 \end{pmatrix} + \alpha_2 \begin{pmatrix} 1 \\ 0 \\ 0 \end{pmatrix}$$

Energy gap in this frame becomes the length of the vector

$$\sqrt{(\omega_\mathrm{m} - k_2)^2 + \alpha_2^2} \to \omega_\mathrm{m} - k_2 + \frac{1}{2} \frac{\alpha_2^2}{\omega_\mathrm{m} - k_2}$$

Relative detuning

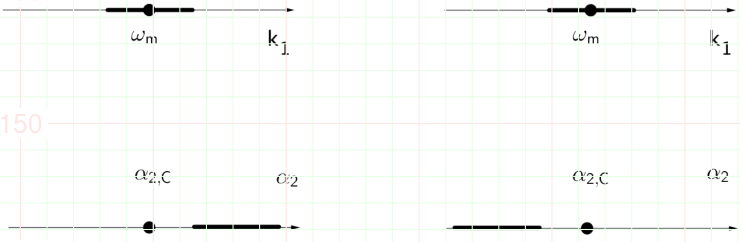
$$D' = \left| \frac{\omega_\mathrm{m} - k_1}{\alpha_1} + \frac{\alpha_2^2}{2\alpha_1(\omega_\mathrm{m} - k_2)} \right|$$

Interferences of Rabi Oscillations

$$D' = \left| \frac{\omega_\mathrm{m} - k_1}{\alpha_1} + \frac{\alpha_2^2}{2\alpha_1(\omega_\mathrm{m} - k_2)} \right|$$

Two driving frequencies k_1, and k_2, with amplitude α_1, and α_2

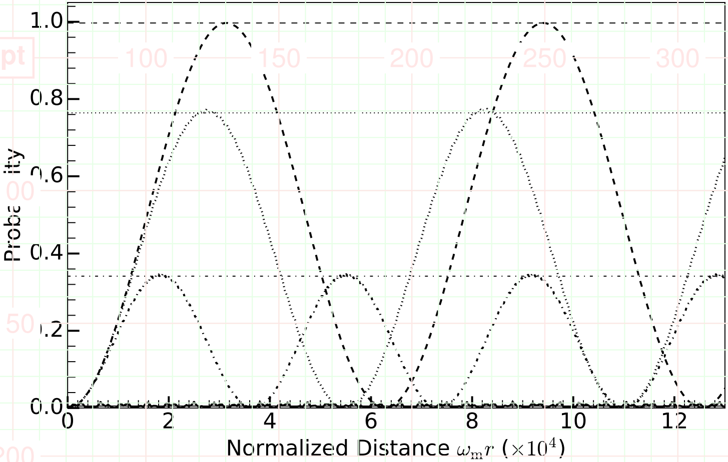
Destruction effect: $k_1 = \omega_\mathrm{m}$, $|\alpha_2| \gg \sqrt{2\omega_\mathrm{m}|\alpha_1(k_2 - \omega_\mathrm{m})|} \equiv \alpha_{2,C}$



Resonance Destruction

Resonance Survival

Interferences of Rabi Oscillations



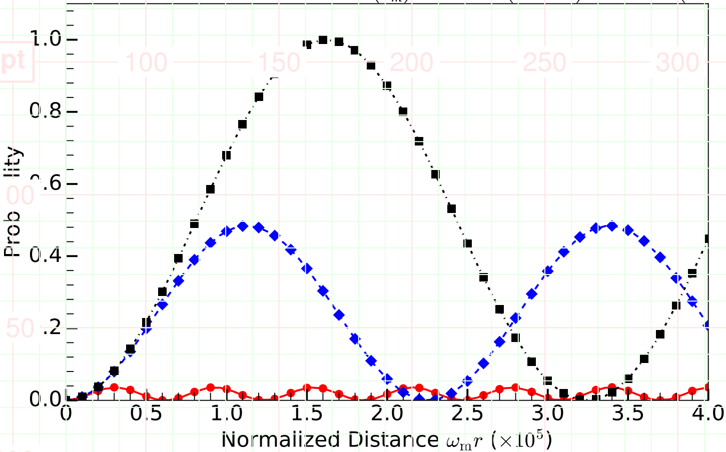
Grid lines: amplitude predicted using $1/(1+D'^2)$

α_2, k_1 values

Dashed	dotted	dash-dotted	solid
$10^{-2}\omega_{\mathrm{m}}, 10\omega_{\mathrm{m}}$	$10^{-2}\omega_{\mathrm{m}}, 10^{-1}\omega_{\mathrm{m}}$	$5.0\times 10^{-2}\omega_{\mathrm{m}}, 10\omega_{\mathrm{m}}$	$5\times 10^{-2}\omega_{\mathrm{m}}, 10^{-1}\omega_{\mathrm{m}}$

Rabi Formula Works

$$\vec{H} = \begin{pmatrix} 0 \\ 0 \\ \omega_m \end{pmatrix} + \frac{A \sin 2\theta_m}{2} \begin{pmatrix} \cos(kx) \\ -\sin(kx) \\ 0 \end{pmatrix} + \frac{A \sin 2\theta_m}{2} \begin{pmatrix} \cos(-kx) \\ -\sin(-kx) \\ 0 \end{pmatrix}$$



Lines: Rabi formula
Dots, diamonds, triangles, and squares are for $k = \omega_m$, $k = (1 - 2\times 10^{-5})\omega_m$, and $k = (1 - 10^{-4})\omega_m$ respectively.

$$\alpha_{2,C} = 2\alpha = 2 \times \frac{A \sin 2\theta_m}{2}$$

Overview

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Jacobi-Anger Expansion
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Summary

Interferences of Rabi Oscillations

We have been making approximations.

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \cos 2\theta_{\mathrm{m}} A\cos(kx)\right)\sigma_3 - \frac{\sin 2\theta_{\mathrm{m}}}{2} A\cos(kx)\sigma_1$$
$$\to -\frac{\omega_m}{2}\sigma_3 - \frac{A\sin 2\theta_{\mathrm{m}}}{2}\cos(kx)\sigma_1$$

We need a better basis.

Rabi Basis

Hamiltonian in Background Matter Basis

$$\mathbf{H} = \frac{1}{2}\left(-\omega_{\mathrm{m}} + \delta\lambda(x)\cos 2\theta_{\mathrm{m}}\right)\boldsymbol{\sigma_3} - \frac{\delta\lambda(x)}{2}\sin\theta_{\mathrm{m}}\boldsymbol{\sigma_1}.$$

A Better Basis

Define Rabi basis $\{|\tilde{\nu}_{\mathrm{L}}\rangle, |\tilde{\nu}_{\mathrm{H}}\rangle\}$ is related to background matter basis $\{|\nu_{\mathrm{L}}\rangle, |\nu_{\mathrm{H}}\rangle\}$ through

$$\begin{pmatrix} |\nu_{\mathrm{L}}\rangle \\ |\nu_{\mathrm{H}}\rangle \end{pmatrix} = \begin{pmatrix} e^{-i\eta(x)} & 0 \\ 0 & e^{i\eta(x)} \end{pmatrix} \begin{pmatrix} |\tilde{\nu}_{\mathrm{L}}\rangle \\ |\tilde{\nu}_{\mathrm{H}}\rangle \end{pmatrix},$$

where

$$\eta(x) - \eta(0) = \frac{\cos 2\theta_{\mathrm{m}}}{2}\int_0^x \delta\lambda(\tau)d\tau.$$

SINGLE FREQUENCY MATTER PROFILE

Matter profile

$$\lambda(x) = \lambda_0 + A\sin(kx),$$

Hamiltonian in new basis

$$\widetilde{\mathbf{H}} = -\frac{\omega_{\mathrm{m}}}{2}\sigma_3 - \frac{\delta\lambda(x)}{2}\sin 2\theta_{\mathrm{m}}\begin{pmatrix} 0 & e^{2i\eta(x)} \\ e^{-2i\eta(x)} & 0 \end{pmatrix} = -\frac{\omega_{\mathrm{m}}}{2}\sigma_3 + \begin{pmatrix} 0 & h \\ h^* & 0 \end{pmatrix}$$

Hamiltonian in New Basis

$$\begin{aligned} h &= -\frac{\delta\lambda(x)}{2}e^{2i\eta(x)} \\ &= \frac{i}{4}\left[\exp\left(i(k+\omega_{\mathrm{m}})x + i\cos 2\theta_{\mathrm{m}}\frac{A}{k}\cos(kx)\right)\right. \\ &\quad \left. - \exp\left(i(-k+\omega_{\mathrm{m}})x + i\cos 2\theta_{\mathrm{m}}\frac{A}{k}\cos(kx)\right)\right] \end{aligned}$$

Single Frequency Matter Profile

Off-diagonal Term in Our System

$$\widetilde{\mathrm{H}} = -\frac{\omega_{\mathrm{m}}}{2}\sigma_3 + \begin{pmatrix} 0 & h \\ h^* & 0 \end{pmatrix}$$

$$h \propto \Bigg[\exp\left(i(k+\omega_{\mathrm{m}})x + i\cos 2\theta_{\mathrm{m}}\frac{A}{k}\cos(kx)\right) - \exp\left(i(-k+\omega_{\mathrm{m}})x + i\cos 2\theta_{\mathrm{m}}\frac{A}{k}\cos(kx)\right)\Bigg]$$

Jacobi-Anger expansion

$$e^{i\beta\cos(kx)} = \sum_{n=-\infty}^{\infty} i^n J_n(\beta) e^{inkx},$$

where $J_n(\beta)$ are Bessel's functions of the first kind.

Single Frequency Matter Profile

Scaled Quantities

Characteristic scale: ω_{m}

- $\hat{A} = A/\omega_{\mathrm{m}}$
- $\hat{k} = k/\omega_{\mathrm{m}}$
- $\hat{x} = \omega_{\mathrm{m}} x$
- $\hat{h} = h/\omega_{\mathrm{m}}$

Single Frequency Matter Profile

Rotation Wave Approximation

e off-diagonal element of Hamiltonian

$$\tilde{\mathbf{H}} = \sum_{n=-\infty}^{\infty} \begin{pmatrix} 0 & \frac{1}{2}\hat{B}_n e^{i(n\hat{k}-1)\hat{x}} \\ \frac{1}{2}\hat{B}_n^* e^{-i(n\hat{k}-1)\hat{x}} & 0 \end{pmatrix}$$

where $\hat{B}_n = -(-i)^n n\hat{k} \tan 2\theta_{\rm m} J_n(\hat{A}\cos 2\theta_{\rm m}/\hat{k})$.

Single Frequency Matter Profile

Transition Probability

$$P^{(n)}_{\mathrm{L}\to\mathrm{H}} = \frac{\left|\hat{B}_n\right|^2}{\left|\hat{B}_n\right|^2 + (n\hat{k}-1)^2} \sin^2\left(\frac{q^{(n)}}{2}x\right),$$

where

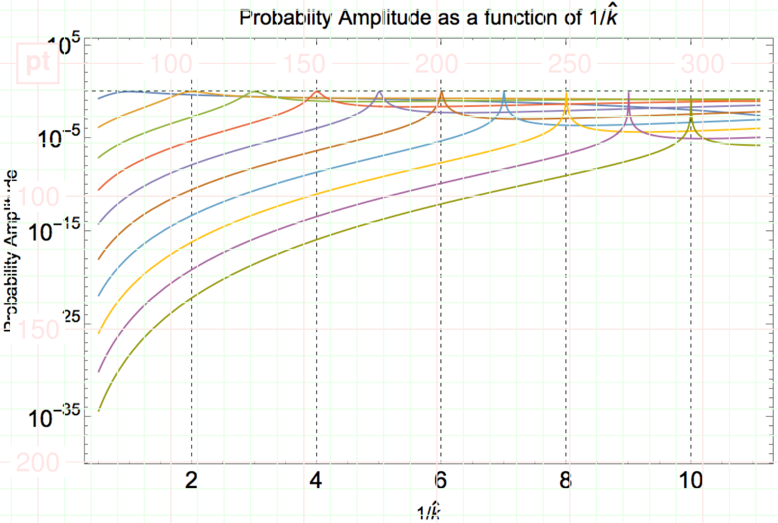
$$q^{(n)} = \sqrt{\left|\Gamma^{(n)}\right|^2 + (n\hat{k}-1)^2}, \quad \text{frequency of oscillations}$$

$$\Gamma^{(n)} = \left|\hat{B}_n\right|, \quad \text{width of resonance } (n\hat{k} \text{ as parameter})$$

Resonance conditions

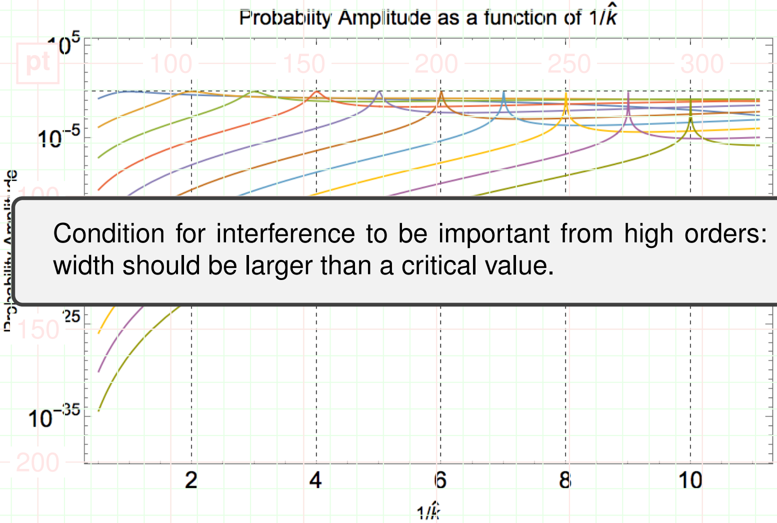
$$\hat{k} \sim \frac{1}{n}$$

Single Frequency Matter Profile



Resonances of different $n = 1/\hat{k}$. Width becomes extremely narrow for high ord

Single Frequency Matter Profile



Condition for interference to be important from high orders: width should be larger than a critical value.

Resonances of different $n = 1/\hat{k}$. Width becomes extremely narrow for high or

Single Frequency Matter Profile Revisited

$k_1 = \omega_{\mathrm{m}}$

n	D	D'_1	$2\pi\omega_{\mathrm{m}}/\Omega_n$
1	0	-	3.2×10^5
-1	10^5	4.8×10^{-6}	3.1
2	1.1×10^9	2.1×10^{-14}	6.3
-2	3.4×10^9	6.9×10^{-15}	2.1

Single Frequency Matter Profile Revisited

$$k_1 = (1 - 2 \times 10^{-5})\omega_\mathrm{m}$$

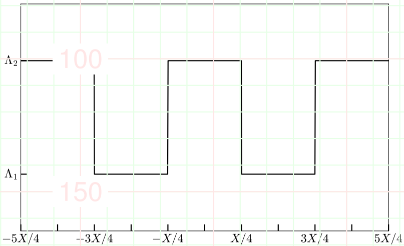
n	D	D_1'	$2\pi\omega_\mathrm{m}/\Omega_n$
1	1	-	2.2×10^5
-1	10^5	1	3.1
2	1.1×10^9	1	6.3
-2	3.4×10^9	1	2.1

SINGLE FREQUENCY MATTER PROFILE REVISITED

$$k_1 = (1 - 10^{-4})\omega_{\mathrm{m}}$$

n	D	D_1'	$2\pi\omega_{\mathrm{m}}/\Omega_n$
1	5.2	-	6.2×10^4
-1	10^5	5.2	3.1
2	1.1×10^9	5.2	6.3
-2	3.4×10^9	5.2	2.1

Castle Wall Matter Profile



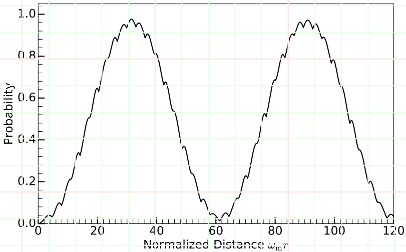
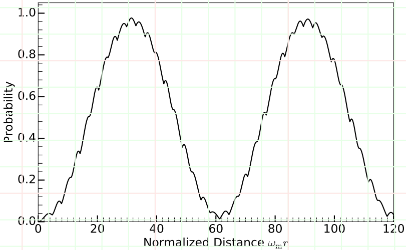


Figure: Castle wall matter profile

Castle Wall Matter Profile

Table: Relative detuning of each frequ

$\{n_1, n_2\}$	D	$D'_{\{1,0\}}$
$\{1, 0\}$	0	-
$\{-1, 0\}$	48	1.0×10^{-2}
$\{0, 1\}$	1.5×10^{2}	1.1×10^{-3}
$\{2, 0\}$	2.4×10^{2}	2.0×10^{-4}



Overview

Introduction

Matter Effect

Stimulated Neutrino Flavor Transitions

Jacobi-Anger Expansion

Summary

SUMMARY

- The fact that neutrino flavor sates are not mass states causes vacuum oscillations.
- MSW resonance happens when matter potential cancels out the vacuum diagonal elements of the Hamiltonian.
- Even matter profile doesn't match MSW requirement, variation in matter profile can cause resonances.
- Single frequency perturbations in matter profile is a combination of many Rabi oscillations.
- Rabi oscillations with two driving fields of different frequencies: large width to destroy the resonance.

Backup Slides

BACKUP SLIDES

Rabi Oscillations

The coupling strength is calculated as

$$\alpha = \langle 1|\mathbf{d} \cdot \mathbf{E}|2\rangle$$

where the electric field is

$$\mathbf{E} = \mathbf{E}_0 \sin(kt).$$

and $\mathbf{d}$ is the dipole moment.

Interferences of Rabi Oscillations

$$-\frac{\omega_0}{2}\sigma_3-\frac{\alpha}{2}\begin{pmatrix}0 & e^{ikt}\\ e^{-ikt} & 0\end{pmatrix}$$

$$\begin{aligned}
\mathbf{H} &= \frac{1}{2}\left(-\omega_{\mathrm{m}} + \cos 2\theta_{\mathrm{m}} A\cos(kx)\right)\sigma_3 - \frac{\sin 2\theta_{\mathrm{m}}}{2} A\cos(kx)\sigma_1\\
&\to -\frac{\omega_m}{2}\sigma_3 - \frac{A\sin 2\theta_{\mathrm{m}}}{2}\cos(kx)\sigma_1\\
&= -\frac{\omega_m}{2}\sigma_3 - \frac{A\sin 2\theta_{\mathrm{m}}}{2}\frac{1}{2}\begin{pmatrix}0 & e^{ikx}\\ e^{-ikx} & 0\end{pmatrix} - \frac{A\sin 2\theta_{\mathrm{m}}}{2}\frac{1}{2}\begin{pmatrix}0 & e^{i(-k)x}\\ e^{-i(-k)x} & 0\end{pmatrix}
\end{aligned}$$

Parameters Used for Vacuum Oscillations

$\theta_{12} = 33.36/180\pi$; $\theta_{13} = 8.66/180\pi$; $\theta_{23} = 40/180 * \pi$; $\delta_{cp} = 0$;
$m_1^2 = 0.01$; $m_2^2 = m_1^2 + 0.000079$; $E = 1\text{MeV}$

Single Frequency Matter Profile

pt 100 150 200 250 300 350

Why Does It Work?

$$J_n(n \operatorname{sech} \alpha) \sim \frac{e^{-n(\alpha - \tanh \alpha)}}{\sqrt{2\pi n \tanh \alpha}}, \quad \text{for large } n$$

$\Rightarrow$

$$\Gamma \propto \hat{B}_n \propto \frac{e^{-n(\alpha - \tanh \alpha)}}{\sqrt{2\pi n \tanh \alpha}}$$

Small perturbation $\Rightarrow$ Small $\hat{A}$ $\Rightarrow$ Large α $\Rightarrow$ Drops fast at large n.

100 150 200 250

Two-frequency Matter Profile

Matter Profile

$$\lambda(x) = \lambda_0 + \delta\lambda(x), \quad \delta\lambda(x) = A_1 \sin(k_1 x) + A_2 \sin(k_2 x).$$

Two-Frequency Matter Profile $\hat{h} = \sum_{n=-\infty}^{\infty} \frac{1}{2}\hat{B}_n e^{i(n\hat{k}-1)\hat{x}}$,

Hamiltonian Off-diagonal Element

Apply Jacobi-Anger expansion,

$$\hat{h} = \sum_{n_1=-\infty}^{\infty} \sum_{n_2=-\infty}^{\infty} \frac{1}{2}\hat{B}_{n_1,n_2}(\hat{k}_1,\hat{k}_2) e^{i(n_1\hat{k}_1+n_2\hat{k}_2-1)\hat{x}},$$

where

$$B_{n_1,n_2}(\hat{k}_1,\hat{k}_2) = -(-i)^{n_1+n_2}(n_1\hat{k}_1 + n_2\hat{k}_2) J_{n_1}\left(\frac{\hat{A}_1 \cos 2\theta_{\mathrm{m}}}{\hat{k}_1}\right) J_{n_2}\left(\frac{\hat{A}_2 \cos 2\theta_{\mathrm{m}}}{\hat{k}_2}\right)$$

Which terms are important?

Two-frequency Matter Profile

Resonance Lines

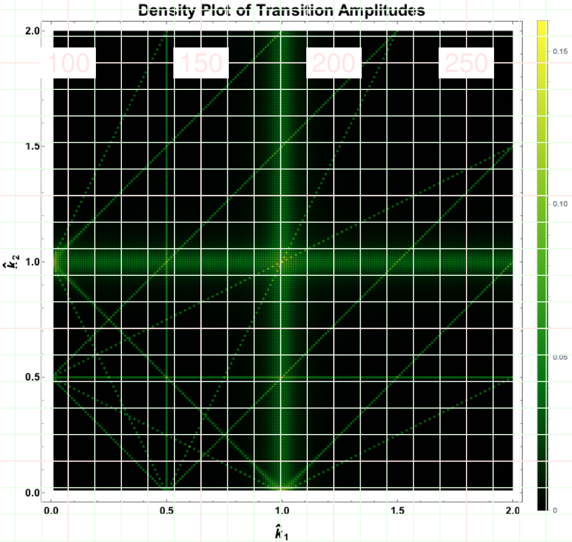
There are still resonances, i.e., (almost) zero phases, on lines

$$n_{1,0}\hat{k}_1 + n_{2,0}\hat{k}_2 - 1 = 0$$

in $\{\hat{k}_1, \hat{k}_2\}$ plane. $\Rightarrow$ Resonance width for each point on resonance lines.

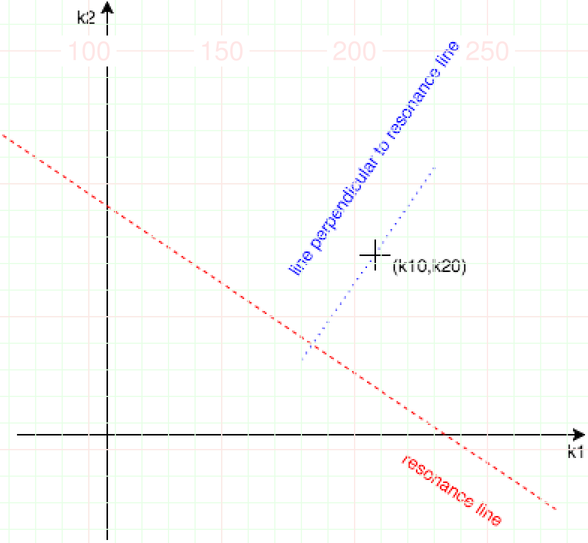
TWO-FREQUENCY MATTER PROFILE

$$\hat{h} = \sum_{n_1}\sum_{n_2} \frac{1}{2}\hat{B}_{n_1,n_2}(\hat{k}_1,\hat{k}_2)\, e^{i(n_1\hat{k}_1+n_2\hat{k}_2-1)\hat{x}},$$



Density plot of transition amplitudes calculated using only one term out of the whole summation in Hamiltonian. $n_1, n_2 \in [-2, 2]$

Two-frequency Matter Profile



Resonance line, distance to resonance, and width

Two-frequency Matter Profile

Width

$$\Gamma_2 = \frac{\hat{B}_{n_1,n_2}(\hat{k}_{1,\text{intercept}}, \hat{k}_{2,\text{intercept}})}{\sqrt{n_1^2 + n_2^2}}.$$

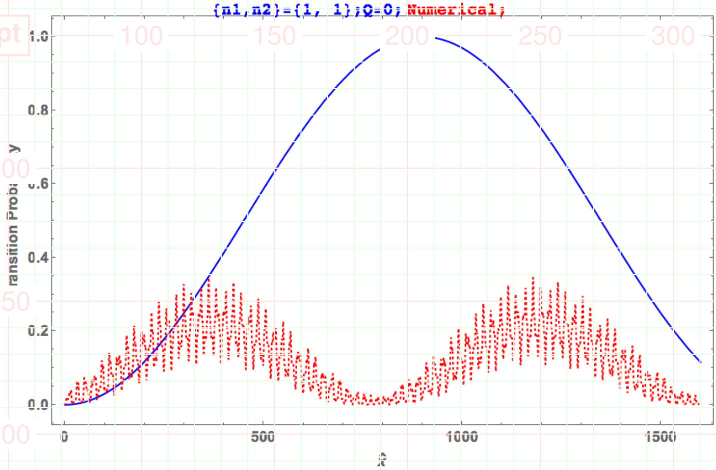
Distance to Resonance Line

$$d = \frac{|n_1\hat{k}_{10} + n_2\hat{k}_{20} - 1|}{\sqrt{n_1^2 + n_2^2}}.$$

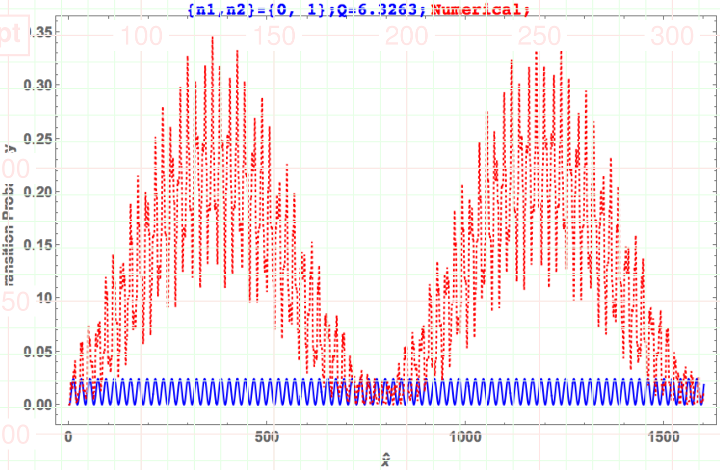
Distance to Resonance Width Ratio

$$Q_2 = \frac{d}{\Gamma_2}.$$

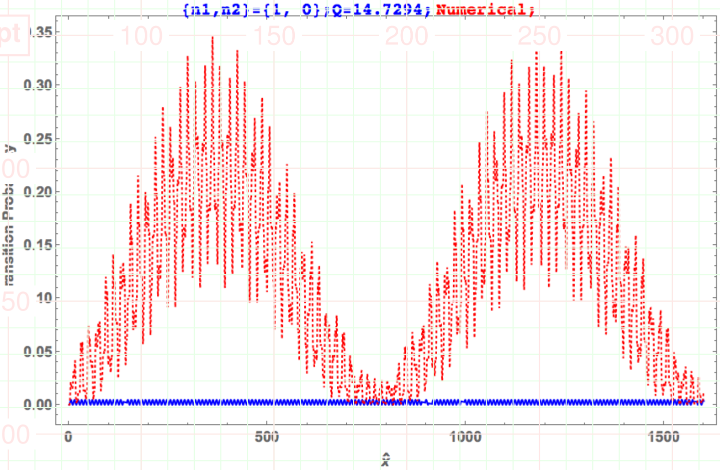
Two-frequency Matter Profile



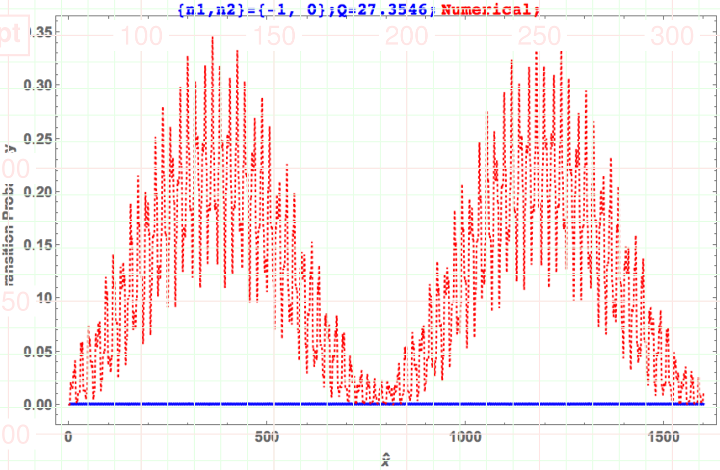
Two-frequency Matter Profile



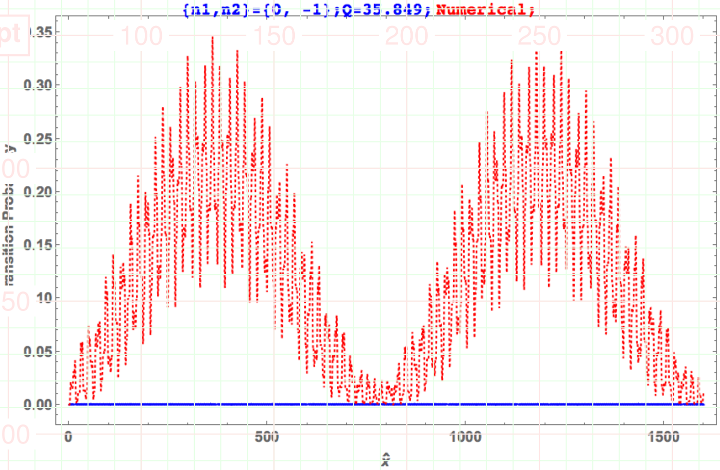
Two-frequency Matter Profile



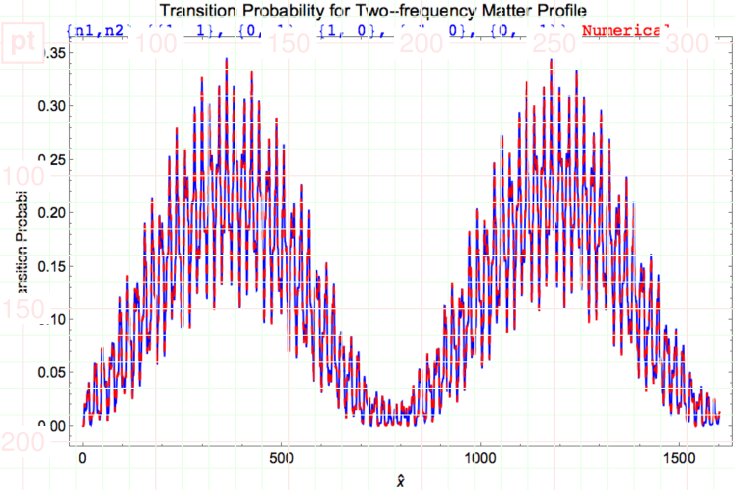
Two-frequency Matter Profile



Two-frequency Matter Profile



Two-frequency Matter Profile



Bessel's Function

$$J_n(\beta) = \sum_{m=0}^{\infty} \frac{(-1)^m}{m!\,\Gamma(m+n+1)} \left(\frac{\beta}{2}\right)^{2m+n}$$

References I