

Internally heated, stratified, compressible convection

Evan H. Anders and Benjamin P. Brown

*Dept. Astrophysical & Planetary Sciences, University of Colorado – Boulder, Boulder, CO 80309, USA and
Laboratory for Atmospheric and Space Physics, Boulder, CO 80303, USA*

An abstract will go here eventually

I. INTRODUCTION

1. Importance of internal heating in natural convective systems
2. Shape of internal heating in something like Sun
3. Studies in Rayleigh-Benard
4. Goals of this paper:
 - (a) Show how to set up internally heated polytropes, with stable layers (or without)
 - (b) Show how the resultant stratification depends on combo of lower flux + internal heating profile choices
 - (c) Show how stable layer affects surface power spectrum

II. EXPERIMENT

We choose a parameter, ϵ , which sets the initial scale of gravity,

$$g \equiv m_{ad} + 1 - \epsilon, \quad (1)$$

where $m_{ad} \equiv (\gamma - 1)^{-1}$, and $\gamma = c_P/c_V = 5/3$. The atmosphere has a constant internal heating term, such that the energy equation looks like

$$\frac{\partial}{\partial t} \left(\left[\rho \frac{|\mathbf{u}|^2}{2} + c_v T + \phi \right] \right) + \nabla \cdot (\mathbf{F}_{\text{conv}} + \mathbf{F}_{\text{cond}}) = \kappa H, \quad (2)$$

where H is an internal heating term. In the initial, static, conductive state, the conductive flux must balance the internal heating term,

$$\nabla \cdot (-\kappa \nabla T_c) = \kappa H, \quad (3)$$

A useful form of the initial temperature gradient in such a system is

$$\nabla T_c = -1 + H(L_z - z). \quad (4)$$

This initial conductive solution is *not* the adiabatic solution to this atmosphere,

$$\nabla T_{ad} = -\frac{g}{c_P} = -1 + \frac{\epsilon}{c_P}. \quad (5)$$

The initial *superadiabatic temperature gradient* is an important quantity, as it determines whether or not the atmosphere is stable to convection. This gradient in these systems is

$$\nabla T_c - \nabla T_{ad} = H(L_z - z) - \frac{\epsilon}{c_P}. \quad (6)$$

When this quantity is zero, the atmosphere is carrying the maximum amount of flux that conductivity can carry. When this value is less than zero, the atmosphere requires convection to carry the flux. Solving for a generalized zero-point, we find that

$$z_{\text{cross}} \equiv L_z - \frac{\epsilon}{H c_P}. \quad (7)$$

In order for the atmosphere to have any region that is unstable to convection, we require that $0 < \epsilon/(Hc_P) < L_z$. In the case that $\epsilon/(Hc_P) < L_z$, the whole atmosphere is unstable to convection. Further, the anticipated length scale of the convecting region of the atmosphere is the full atmosphere above z_{cross} , or

$$d_{\text{conv}} = L_z - z_{\text{cross}} = \frac{\epsilon}{Hc_P} \quad (8)$$

I think there are two natural control parameters of this system:

1. $f \equiv z_{\text{cross}}/L_z$, the fraction of the atmosphere that is stable ($f = 0$ means the full atmosphere convects, $f = 0.5$ means that the upper half convects, etc). This *should* be the parameter that determines the depth of the stable layer.
2. ϵ , the deviation of the initial temperature gradient (at the top of the atmosphere) from the adiabatic temperature gradient. This *should* be the parameter that determines the depth of the stable layer.

By specifying these two parameters, we retrieve the magnitude of the internal heating, $f = 1 - \epsilon/(Hc_PL_z)$

$$H = \frac{\epsilon}{L_z c_P (1 - f)}. \quad (9)$$

There are still three parameters in these systems that need to be figured out:

1. Density stratification
2. thermal diffusivity
3. viscous diffusivity

One of these parameter is free,

$$\text{Pr} = \frac{\nu}{\chi}. \quad (10)$$

We're not going to mess with that.

There will be some type of Rayleigh number in these systems,

$$\text{Ra} = \frac{\text{stuff that comes from the atmosphere}}{\nu\chi}, \quad (11)$$

but I need to think a little more and spend some time with the equations to figure out what this is.

And then as for the stratification....well, I need to figure out what n_ρ I'm specifying. The n_ρ of the corresponding adiabatic polytrope?

III. RESULTS & DISCUSSION