

# Controlling the rotational constraint in stratified, compressible convection

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## ABSTRACT

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### 1. INTRODUCTION

Convection under the influence of rotation has been studied in great detail in recent decades. In the incompressible boussinesq case, numerous authors have studied rotating convection in both laboratory and numerical settings. The behavior of the heat transport in these systems has been the focus of many of these studies, particularly the difference of behavior of heat transport in rotationally-constrained and unconstrained regimes (King et al. 2009; Zhong et al. 2009; Stevens et al. 2009; Julien et al. 2012). Studies of both laboratory (cite, cite) and numerical (cite, cite) experiments coincide well in these regimes.

More complicated experiments of rotational convection are often conducted using numerical tools in spherical geometries. Often these studies aim to gain insight into the solar dynamo (Glatzmaier & Gilman 1982; Busse 2002; Brown et al. 2008, 2010, 2011; Augustson et al. 2012; Guerrero et al. 2013; Käpylä et al. 2014). Numerous interesting phenomenological discoveries have been made in such studies, but they often employ very different degrees of rotational constraint, but their basic system behavior often differs significantly from the sun. For example, dynamo simulations often produce antisolar differential rotation profiles at nominally solar values of rotational constraint.

In recent years, helioseismic imaging of flows in the Sun have suggested that power in large-scale convective motions is much lower than convective simulations and mixing length theory predict (Hanasoge et al. 2012;

Greer et al. 2015). This suggests that either convection is not driven deep in the solar convection zone or its motions are masked before they reach the surface by some process. A number of possible mechanisms have been suggested as culprits for this behavior, including entropy rain (Brandenburg 2016). One other possibility is that the interior convection zone is rotationally constrained, and that this behavior is reducing low wavenumber power at the surface. The recent work of Featherstone & Hindman (2016) suggests that this is a strong possibility, and the observations of Greer et al. (2016) suggest that flows in the deep solar convection zone are rotationally constrained. The degree of rotational constraint under which convective flows occur can greatly change the character of the resultant situation, and so in the case of astrophysical contexts (solar studies such as those mentioned above and also planetary studies such as e.g., Soderlund et al. (2015)), it is important to study convection in the same rotational regime as the astrophysical object of interest in order to meaningfully understand results.

One difficulty in studying rotational convection is that it is often unclear from input parameters whether or not the resulting convective state will be rotationally constrained. In Anders & Brown (2017) (hereafter AB17), we studied non-rotating, hydrodynamic, compressible convection. We showed how the evolved Reynolds number, Peclet number (Re and Pe, two measures of turbulence in the evolved state), and Mach number (Ma, the ratio of flow speed to the local sound speed) of the convective solution can be specified through a properly constructed reference atmosphere. Upon the inclusion of rotation, a fourth dynamical measure of the solution becomes meaningful: the Rossby number (Ro, the ratio of advective dynamics to rotational constraint). Low Ro

flows are rotationally constrained, while high  $Ro$  flows are not. While the literature contains a wealth of information regarding how the magnitude of rotation affects the heat transport characteristics of convection, we find no work which simply links the rotational constraint of evolved solutions with input parameters.

In this work, we extend the study of AB1717 to rotationally-influenced,  $f$ -plane polytropic atmospheres, as have been previously studied by e.g., Brummell et al. (1996, 1998); Calkins et al. (2015). Our goal is to determine how the input parameters which we studied previously couple with a new input parameter, the Taylor number ( $Ta$ , Julien et al. (1996)), which sets the magnitude of the rotational vector. In section 2, we describe our atmosphere, numerical experiment, and paths through parameter space. In section 3, we present the results of our experiments and we offer concluding remarks.

## 2. EXPERIMENT

We study fully compressible, stratified convection as we previously did in AB17, with the inclusion of rotation. We study an ideal gas whose equation of state is  $P = R\rho T$  and with an adiabatic index  $\gamma = 5/3$ . We nondimensionalize the atmosphere such that  $R = 1$  and  $P = \rho = T = 1$  at the top of the domain. The initial stratification is polytropic, such that

$$\begin{aligned}\rho_0(z) &= (1 + L_z - z)^m, \\ T_0(z) &= (1 + L_z - z),\end{aligned}\tag{1}$$

where  $m$  is the polytropic index,  $z$  increases upwards in the range  $z = [0, L_z]$ , and  $L_z \equiv e^{n_\rho/m} - 1$  is the depth of the atmosphere, where  $n_\rho$  specifies the number of density scale heights that the atmosphere spans. We specify the instability of the atmosphere through the superadiabatic excess,  $\epsilon = m - m_{ad}$ , where  $m_{ad} = (\gamma - 1)^{-1}$  is the adiabatic polytropic index, and  $\epsilon$  controls the Mach number of the flows (Anders & Brown 2017). The domain is a 3D cartesian box whose horizontal extent is in the range  $x, y = [0, AL_z]$ , where  $A$  is the aspect ratio of the domain. As has been studied previously by e.g., Julien et al. (1996); Brummell et al. (1996), we study a domain in which the gravity and rotational vector are antiparallel,  $\mathbf{g} = -g\hat{z}$ , and  $\mathbf{\Omega} = \Omega\hat{z}$ .

We evolve the velocity ( $\mathbf{u}$ ), temperature ( $T$ ), and log density ( $\ln \rho$ ) according to the Fully Compressible

Navier-Stokes equations,

$$\frac{\partial \ln \rho}{\partial t} + \nabla \cdot \mathbf{u} = -\mathbf{u} \cdot \nabla \ln \rho,\tag{2}$$

$$\begin{aligned}\frac{\partial \mathbf{u}}{\partial t} + 2\mathbf{\Omega} \times \mathbf{u} + \nabla T - \nu \nabla \cdot \bar{\bar{\boldsymbol{\sigma}}} - \bar{\bar{\boldsymbol{\sigma}}} \cdot \nabla \nu = \\ -\mathbf{u} \cdot \nabla \mathbf{u} - T \nabla \ln \rho + \mathbf{g} + \nu \bar{\bar{\boldsymbol{\sigma}}} \cdot \nabla \ln \rho,\end{aligned}\tag{3}$$

$$\begin{aligned}\frac{\partial T}{\partial t} - \frac{1}{c_V} (\chi \nabla^2 T + \nabla T \cdot \nabla \chi) = \\ -\mathbf{u} \cdot \nabla T - (\gamma - 1) T \nabla \cdot \mathbf{u} \\ + \frac{1}{c_V} (\chi \nabla T \cdot \nabla \ln \rho + \nu [\bar{\bar{\boldsymbol{\sigma}}} \cdot \nabla] \cdot \mathbf{u}),\end{aligned}\tag{4}$$

with the viscous stress tensor given by

$$\sigma_{ij} \equiv \left( \frac{\partial u_i}{\partial x_j} + \frac{\partial u_j}{\partial x_i} - \frac{2}{3} \delta_{ij} \nabla \cdot \mathbf{u} \right).\tag{5}$$

The kinematic viscosity ( $\nu$ ), thermal diffusivity ( $\chi$ ), and strength of rotation ( $\Omega$ ) are set at the top of the domain by the Rayleigh number ( $Ra$ ), Prandtl number ( $Pr$ ), and Taylor number ( $Ta$ ),

$$Ra = \frac{gL_z^3 \Delta S / c_P}{\nu \chi}, \quad Pr = \frac{\nu}{\chi}, \quad Ta = \left( \frac{2\Omega L_z^2}{\nu} \right)^2,\tag{6}$$

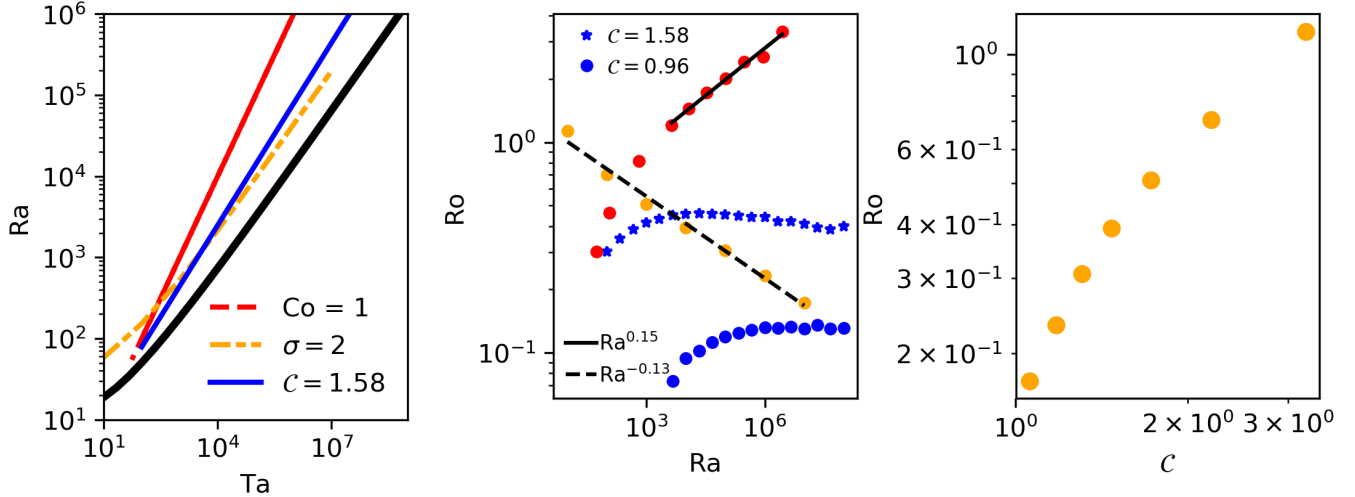
where  $\Delta S = \epsilon n_\rho / m$  is the specific entropy difference between  $z = 0$  and  $z = L_z$ , and the specific heat at constant pressure is  $c_P$ .

As  $Ta$  increases, the critical value of  $Ra$  at which convection onsets,  $Ra_{crit}$ , also increases. The linked nature of these crucial control parameters makes it difficult to predict the rotational constraint of the evolved fluid flows for a given set of input parameters. In this work, we will explore three paths through  $Ra$ - $Ta$  space:

$$Ra = \begin{cases} \mathcal{S} Ra_{crit}(Ta), & \text{(I)} \\ Co^2 Pr Ta, & \text{(II)} \\ \mathcal{P}_{Ro}^2 Pr Ta^{3/4} & \text{(III).} \end{cases}\tag{7}$$

Paths on constraint I are at constant supercriticality,  $\mathcal{S}$ . Paths on constraint II are at a constant value of the classic ‘‘Convective Rossby number’’ ( $Co$ ), which has been used frequently over the past two decades, and is intended to predict the rotational constraint of the evolved solution (Julien et al. 1996; Brummell et al. 1996). Paths on constraint III set constant a ratio which we call the ‘‘Predictive Rossby Number’’ ( $\mathcal{P}_{Ro}$ ). These paths follow contours of  $Ta^{3/4}$ , and to our knowledge have not been well-explored in the literature. An example of each of these paths can be seen in Fig. 1a.

In this work, our goal is to study the nature of rotational constraint along each path defined in Eqn. 7.



**Figure 1.** (a) The critical Rayleigh number, as a function of the Taylor number, is plotted as a solid black line. Paths of constant Convective Rossby Number (red dashed line), constant supercriticality (orange dashed line), and  $\mathcal{P}_{Ro}$  (blue solid line) are shown through parameter space. (b) Evolved Rossby number is plotted vs. Taylor number along multiple constant  $\mathcal{P}_{Ro}$  paths, such as the solid blue line in (a). After a sharp increase at low  $Ta$ , the evolved Rossby number flattens out and stays nearly constant across orders of magnitude of  $Ta$ .

We will quantify rotational constraint using the Rossby number,

$$Ro = \frac{|\nabla \times \mathbf{u}|}{2\Omega}. \quad (8)$$

We evolve our simulations using the Dedalus<sup>1</sup> pseudospectral framework, and our numerical methods are identical to those presented in AB17.

### 3. RESULTS & DISCUSSION

In Fig. 1a, we display the value of the critical Rayleigh number ( $Ra_{crit}$ ) as a function of the Taylor number ( $Ta$ ). We found these critical values through the use of a linear instability analysis. We plot sample paths through this parameter space at  $\sigma = 1$ ,  $Co = 1$ , and  $C = 0.957$ , as defined in Eqn. 7. In Fig. 1b, we show that along paths of constant  $\sigma$ , the Rossby number, as defined in Eqn. 8, decreases weakly as a function of  $Ra$  and  $Ta$ . Along paths of constant  $Co$ ,  $Ro$  increases sharply with increasing  $Ra$  and  $Ta$ . However, along paths of constant  $C$ , after an increase at low  $Ra$ , the evolved value of  $Ro$  levels off and stays roughly constant as a function of  $Ra$ . By interpolating between the slopes of constant  $\sigma$  and constant  $Co$  lines, we arrived at the scaling of  $Ra \propto Ta^{3/4}$  along constant  $C$  lines, and holding  $C$  constant seems to predict  $Ro$  relatively well from our experimental trials here. In Fig. 1c, we show how the Predictive Rossby Number,  $C$ , relates to the evolved Rossby number for a series of runs which are all at the same supercriticality.

We define the Nusselt number ( $Nu$ ) as we did previously in AB17,

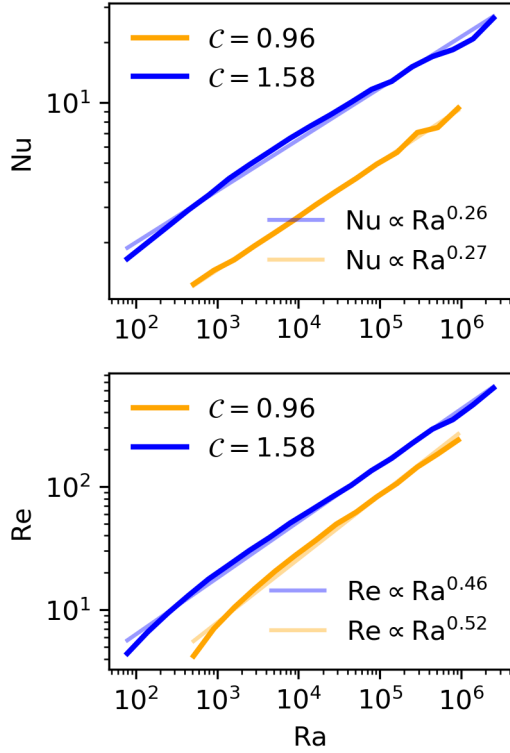
$$Nu \equiv \frac{\langle F_{conv,z} + F_{cond,z} - F_A \rangle}{\langle F_{cond,z} - F_A \rangle}, \quad (9)$$

where  $F_{conv,z}$  is the vertical convective flux,  $F_{cond,z}$  is the vertical conductive flux, and  $F_A$  is the adiabatic conductive flux. We refer the reader to AB17 for a more thorough definition of these fluxes. In Fig. 2a, we show the scaling of the Nusselt number as a function of  $Ra$  at fixed  $C$ . We find that when  $C$  is held constant and  $Ra$  is increased, we find a scaling of  $Nu \propto Ra^{2/7}$  [note: insert true scaling], reminiscent of classical scaling laws in non-rotating theory as  $Ra$  increases. Furthermore, in Fig. 2b, we plot the RMS Reynold's number ( $Re = |u|L_z/\nu$ ) as a function of  $Ra$ , and find that  $Re \propto Ra^{1/2}$ , which is reminiscent of how the Reynold's number scales with increasing  $Ra$  in non-rotating convection.

In Fig. 3, we show sample snapshots of the convective flows in our domain. I'll need to talk about them once I get there.

In Fig. 4, we show time- and horizontally-averaged profiles of the rossby number and the entropy gradient. As  $Ra$  increases at a constant value of  $C$ , both the entropy and rotational boundary layers shrink. We measure the upper boundary layer thicknesses of both profiles. We measure the boundary layer thickness of the Rossby number as the distance from the upper boundary to the point at which  $Ro$  is maximized. We measure the boundary layer thickness of the entropy profile by fitting a line to the upper 10 points in the domain, and

<sup>1</sup> <http://dedalus-project.org/>



**Figure 2.** (a) Evolved Nusselt number vs. Rayleigh number /  $Ra_{crit}$  along constant  $\mathcal{C}$  paths. A classic scaling law of  $Nu \propto Ra^{2/7}$  is observed. (b) Evolved Reynolds number vs. Rayleigh number /  $Ra_{crit}$  along constant  $\mathcal{C}$  paths. A classic scaling law of  $Re \propto Ra^{1/2}$  is observed. These laws are reminiscent of standard scaling laws of  $Re$  and  $Nu$  in non-rotational convection (SOURCES SOURCES SOURCES). This suggests that at fixed Rossby number on a constant  $\mathcal{C}$  path (Fig. 1), varying the Rayleigh number affects the evolved dynamics in a manner similar to a nonrotating fluid.

**Figure 3.** Here’s some pretty plots, we’ll figure out what we plot later.

**Figure 4.** (a) Horizontally-averaged profiles of the Rossby number are shown vs.  $z$  for a constant  $\mathcal{C} = X$ . (b) Horizontally-averaged profiles of the entropy gradient are shown vs.  $z$  for a constant  $\mathcal{C} = X$ . (c) Vorticity boundary layer thickness normalized by entropy boundary layer thickness as a function of  $Ta/Ta_{crit}$  for multiple  $\mathcal{C}$  paths. When this measure is  $\gg 1$ , we expect the flows to be buoyancy dominated, when it is  $\ll 1$ , we expect the flows to be rotationally dominated, and when it is  $\sim 1$ , we anticipate that both effects are very important.

assuming that the boundary layer extends to the point at which that line crosses through zero. We plot  $\delta_S/\delta_{Ro}$  as a function of  $Ra$  for our  $\mathcal{C}$  paths in Fig. 4b. We find that the ratio of these boundary layers is relatively constant as a function of  $Ra$ , which implies that the rotational constraint across  $Ra$  is roughly constant. [NOTE: NEED TO TALK ABOUT MAGNITUDE OF RO IN BOUNDARY LAYER]

#### 4. DISCUSSION

In this letter, we studied rotating, stratified, compressible convection at low Mach number. We studied traditional paths through  $Ra$ - $Ta$  space, as well as a new path, in which the Predictive Rossby number  $\mathcal{C}$  is held constant, and  $Ra = \mathcal{C}Ta^{3/4}$ . Along these paths, we find that the evolved rotational constraint, as measured by the Rossby number ( $Ro$ ), is roughly constant as  $Ra$  increases. Furthermore, the heat transport, measured by the Nusselt number ( $Nu$ ), and the level of turbulence, as measured by the Reynolds number ( $Re$ ), vary according to traditional scaling laws as  $Ra$  increases. Together, these phenomena suggest that experimenters in stratified convection can specify the degree of rotational constraint in their evolved solutions by choosing  $\mathcal{C}$ , and then increase  $Ra$  in a manner analogous to unrotating convection (e.g., AB17) to increase the efficiency and turbulent nature of the convective dynamics in question.

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