late Kab, we must find Varo, But dean space we have Varo = Va (rro) = r Varo + ro Var Varo = r (gab - raro) + result follows from = dr² = d(x²+y²+z²) = 2 x dx + ? y dy + ? z dz ar = r ex + r ex + r ex = r = ra r is already projected on both nd the result follows immediately. rectors ûa and ka are coveriently nt, and therefore commutate with hives. The projection Pab therefore as well, and hat satisfies the field on I hat = I hab. To check that it transverse traceless gavge, we calculate (Pampo - z Pab pmn) = 0 = ka (pmpo - z Pat pmn), rach Transverse = Rappo - z Pab pmn) = pmn - z · z · pmn = 0

projector the second, because it projects into a two-dimensional space.

b) Here, we note first that

Meanwhile, the trace part of Amn is clearly killed in the tensor projection, so

[Pam Pbm - 12 Pab Pmn]=

= [Pai Pbi - z Pas Pis][ŋim ŋin - z ŋij ŋmn]

 $= \left[\left(\bar{\eta}_{a} - \hat{\kappa}_{a} \hat{\kappa}^{i} \right) \left(\bar{\eta}_{b} - \hat{\kappa}_{b} \hat{\kappa}^{j} \right) \right]$

- = () ab - Ra Rb) () [) - R' R')] [] : "] ; " - = 37; 7 mm]

 $= [\tilde{\eta}_{a} i \tilde{\eta}_{j}]^{n} - \hat{\kappa}_{a} \hat{\kappa}^{i} \tilde{\eta}_{b}^{j} - \hat{\kappa}_{b} \hat{\kappa}^{j} \tilde{\eta}_{a}^{i} + \hat{\kappa}_{a} \hat{\kappa}_{b} \hat{\kappa}^{i} \hat{\kappa}^{j}$

+ をしずるからんのんのんがなり「ガーカラッカーラブンガーリ

The result follows by combining the last two terms in the first bracket and noting that Amn is symmetric.

3 The geodesic deviation equation derived in class for linearized fields gives

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Here, aa is the first-order relative acceleration of a pair of test masses that are at vest in the Newtonian frame of the Minkowski background, separated by the spatial displacement 36. The right side is already first-order because of the metric perturbation. so we can use the zeroth-order approximant for 36, which is constant, since the metric perturbation perturbation is harmonic,

This shows that the relative motion is also harmonic, with frequency w, a part from two constants of integration corresponding to the initial relative displacement and velocity of the particles. The former is 3°, while the latter vanishes.

We can determine the radius of motion by computing the norm of aa:

aa = z hab 3 = - zw2 A Re [eabe-iwt] 3 b

= - = w2 Ad Re[e-iwt. /2] (xaxb-yayb+zix(ayb) xb

= - INT WZ Ad Re[e-iwt (xa + i ja)]

= - 1 ZVZ WZAd (Racos wt + Jasin wt)

 $= ||a|| = \frac{1}{2\sqrt{z}} w^2 A d = w^2 r = v = \frac{A d}{2\sqrt{z}}$

L

Thus, the radius of the motion scales with the amplitude A of the vave, which of course makes sense.

Circular Motion

Background

displace ment

4 a) The source energy density is

$$P(t,x,y,z) = m \delta(z) \left[\delta(x-a\cos \omega t) \delta(y-a\sin \omega t) + \delta(x+a\cos \omega t) \delta(y+a\sin \omega t) \right]$$

Therefore, the temporal metric perturbation is

$$h^{\circ\circ} \stackrel{\sim}{=} \frac{4}{r} \int t^{\circ\circ}(y) d^3y = \frac{8m}{r}$$

The temporal spatial components vanish;

because of the symmetry of the source. Finally the spatial components are

$$= \frac{2}{r} \frac{d^2}{dt^2} m \begin{vmatrix} 2a^2 \cos wt & 2a^2 \sin wt \cos wt & 0 \\ 2a^2 \sin wt \cos wt & 2a^2 \sin^2 wt & 0 \\ 0 & 0 & 0 \end{vmatrix}$$

$$= \frac{z m a^2}{r} \frac{d^2}{dt^2} \begin{cases} 1 + \cos 2\omega t & \sin 2\omega t & 0 \\ \sin 2\omega t & 1 - \cos 2\omega t & 0 \\ 0 & 0 & 0 \end{cases}$$