

Black Hole and Jets

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1 Lecture 1 — Jet Acceleration

Lets start with an overview of where do we see black holes in the universe. There are supermassive black holes with $M \sim 10^{6-10} M_{\odot}$, which sit in the center of galaxies. There are also stellar mass black holes we can see, which could be in binaries, or associated with Gamma-ray bursts or mergers with neutron stars. Recently there is also evidence of $M \sim 10^{2-5} M_{\odot}$ intermediate mass black holes which might be ultra-luminous X-ray sources. All of these black holes produces jets. We don't really need an event horizon to produce a jet though. We can have jets from stars, neutron stars and WDs as well. There must be some universal properties of jet production that do not depend too much on the exact dynamics of the central object.

How do we observe jets? There is a famous jet 3C279 which is traveling towards us at $v \approx 0.997c$. Its sky location appears to be moving superluminally, but it is an artifact of it moving towards us. We observe these jets at multiple wave bands, from radio, IR/optical, to X-rays and Gamma rays. We can also potentially detect cosmic rays, neutrinos, and gravitational waves from them.

Jets can also affect the galaxies and clusters around them. Examples can be the Perseus cluster, M87, and MS0735.6 systems. The jets extract a large fraction of the black hole energy and throw them into the ambient environment.

The masses of the central BHs and their host galaxies rotation curve are correlated. This indicates a way for the BH to know about the galaxy around it. The AGN radio luminosity also shows radio quiet and radio loud dichotomy, which seems to suggest that there is at least an additional hidden parameter about the jet other than M and \dot{M} . Possibilities can be magnetic flux difference, effect of the ambient medium instead of the BH itself, or it could be BH spin.

Different jets can also have different morphology, for example the Cygnus A galaxy and M87 galaxy, which host similar mass black holes, but have very different image of their jets (Fanaroff & Riley 1974).

In April there will be more observing tools available in the form of the Event Horizon Telescope (EHT), which can directly look at the vicinity of the black hole in the multiple of horizon size.

Lets look at a simple illustration of how jet is produced. Consider a rotating central object and a ceiling above it which is perfectly conducting, representing the ambient medium. Starting with a straight field line, since it is frozen into the ambient medium, as the central object rotates the magnetic field will be wound up, and when there is enough pressure due to the development of magnetic pressure it will push away and into the interstellar medium.

What powers the outflow? Rotation plus large scale magnetic flux leads to jets. In reality, the black hole is very simple. It has only two hairs: mass and spin. It knows nothing about the magnetic flux. However, external accretion can provide magnetic flux threading the black hole event horizon.

I fell asleep...

1.1 Jet Acceleration

Consider a neutron star which is not rotating, with field lines sticking out from the surface. If we set it rotating, the field lines will start to lag behind. The field line perturbation will travel along the field at Alfven speed. We assume the jets are massless (force free) for simplicity. In this limit the Alfven velocity will be speed of light $v_A = c$. We can find B_ϕ by

$$B_\phi = -\frac{v_\phi}{c} B_r = -\frac{\Omega R}{c} B_r \quad (1.1)$$

The electric field is exactly the same as the toroidal magnetic field

$$E = \left| -\frac{\mathbf{v}}{c} \times \mathbf{B} \right| = \frac{\Omega R}{c} B_r \quad (1.2)$$

Therefore the drift velocity of the plasma is can be calculated from E cross B , and find the Lorentz factor

$$\frac{v}{c} = \frac{E}{B}, \quad \gamma^2 = \frac{1}{1 - v^2/c^2} = 1 + (\Omega R/c)^2 \quad (1.3)$$

This means that near the stellar surface, the Lorentz factor will be around 1, but outside the light cylinder the Lorentz factor scales linearly with distance.

However this kind of acceleration can't continue forever, because the outflow has mass attached to it. Lets look at the conserved quantities along jets, which are ratios of conserved fluxes. One conserved quantity is

$$\eta = \frac{F_M}{F_B} = \frac{\gamma \rho v_p}{B_p} = \text{const} \quad (1.4)$$

which is the ratio of the magnetic flux to the mass flux. The second is the ratio of the energy flux over mass flux

$$\mu = \frac{F_E}{F_M} = \gamma \frac{F_{EM}}{F_{KE}} + \gamma \quad (1.5)$$

The ratio F_{EM}/F_{KE} is called the magnetization $\sigma = b^2/4\pi\rho c^2$. This means that μ sets the maximum Lorentz factor. Also σ sets the speed of the fast waves $\gamma_F = \sigma^{1/2}$. In force-free the fast waves travel at speed of light. In reality when the jet Lorentz factor catches up with this fast wave, force-free approximation breaks down

$$\gamma = \gamma_F = \sigma^{1/2} = (\mu/\gamma)^{1/2}, \quad \gamma = \mu^{1/3} \ll \mu \quad (1.6)$$

Because acceleration cease to be linear after a certain radius, we can't really get to the limit $\gamma \sim \mu$, but way less than that.

In 2D γ doesn't saturate exactly at $\mu^{1/3}$, but at

$$\gamma \propto \mu^{1/3} \log^{1/3} r \quad (1.7)$$

Our key assumption above was that $B_\phi^2 - E^2 \ll B_r^2$. Because of this, toroidal magnetic pressure contribution breaks force balance, bending the field lines, and fast jets can't make sharp turns. In the opposite limit we ignore B_r , we have

$$\gamma^2 = \frac{B_\phi^2}{B_\phi^2 - E^2} \quad (1.8)$$

If we consider the force in more detail, one force is the pressure gradient $F_m = -\nabla p_m = p_m/R$. The other is the centrifugal force $F_c = \epsilon_m \gamma^2 / R_c$. If we want these forces to be equal, then

$$\gamma = \left(\frac{R_c}{R} \right)^{1/2} \quad (1.9)$$

Now we can combine both limits

$$\frac{1}{\gamma^2} = \frac{1}{\gamma_1^2} + \frac{1}{\gamma_2^2} \quad (1.10)$$

where γ_1 and γ_2 are the two limits we got in the two different extremes.

Now lets look back at the fluxes. The electromagnetic energy flux is

$$F_{EM} = \frac{cEB_\phi}{4\pi} = \frac{\omega^2 R^2 B_\phi^2}{4\pi c} \quad (1.11)$$

we can rewrite μ as

$$\frac{\Omega^2}{4\pi^2 \eta c} \pi B_p R^2 + \gamma = \frac{\mu}{\Phi} \pi B_p R^2 + \gamma \quad (1.12)$$

where Φ is the magnetic flux $\Phi = (\pi B_p R^2)_F$. Therefore there will be no acceleration if the magnetic flux remains at large distances:

$$\frac{\gamma}{\mu} = 1 - \frac{\pi B_p R^2}{\Phi} \quad (1.13)$$

How do hydro accelerate fluids? We need to make a nozzle, and in MHD we need something similar. We need magnetic flux bunching toward jet axis.

Outside the fast point γ_F all signals travel inside the cone with opening angle

$$\xi = \frac{\gamma_F}{\gamma} \approx \frac{\sigma^{1/2}}{\gamma} \quad (1.14)$$

For communication across jet requires that $\theta \lesssim \xi$. Thus the acceleration limit can be relaxed

$$\gamma \lesssim \frac{\mu^{1/3}}{\theta^{2/3}} \quad (1.15)$$

Therefore jets accelerate much better near the axis!

However, most jets are collimated and are not monopolar. How do collimated jets accelerate? Collimated jets by its geometry has to keep in causal contact. All signals travel inside the Mach cone

$$\xi \leq \frac{\gamma_F}{\gamma} = \frac{\sigma^{1/2}}{\gamma} \quad (1.16)$$

We expect that $\gamma\theta \lesssim \sigma^{1/2} \lesssim 1$. In AGN jets this is observed, but in GRBs we have $\gamma\theta \sim 10 - 100$. What does this mean? One possibility is that GRB jets are unconfined after a certain radius which could be the surface of a star. When the magnetic field lines are unconfined they can expand and create bigger pressure gradient.

2 Lecture 2 — Disk Jet Connection

We are interested what sets the disk-jet connection and what comes out of it. We want to understand as much as possible what is happening at the central engine. Important questions include how is $\dot{M}c^2$ related to jet acceleration, how do disks and jets emit, what can we learn from the spectra and images on small/large scales. We also want to explore what is a health jet diet, i.e. how to kill a jet.

When are jets produced in real systems? Jets are produced TDEs, ultra-luminous X-ray sources, and GRBs. In these events $\lambda = L/L_{\text{edd}} \sim 1$. In these systems the disk is thick and temperature is high. A lot of the radiation is absorbed by the disk, and only a fraction comes out.

They are also produced in Quasars, X-ray binaries, and TDEs, with $\lambda \sim 0.01$. The disk retain its ability to cool and the disk can be razor thin. α viscosity models is good for this kind of disks. The photons can escape easily from the disk.

If we reduce λ even lower to 10^{-5} , the flow becomes dilute, and we again have a thick disk because they can't radiate photons efficiently. This applies to low-luminosity AGN and X-ray binaries.

Fender+ (2004) plotted optical hardness together with intensity (turtlehead diagram). In the low luminosity extreme the radiation spectrum is very flat as a power law, with very hard slope. With higher accretion rate it becomes brighter and at some point it start to transition to a soft spectrum. When the flow comes towards the soft state it will cross the so-called jet line, the black hole jet start to emit blobs, and we call this transient jets. In the soft state there is no jet, and the disk is razor thin and optically thick. Fender+ (2010) looked at the correlation between BH spin and the outflow and they didn't find any. However there is a correlation between the BH spin and transient jets.

How do we even measure the BH spin? There are two major methods, both relying on measuring the “hole” in the center of the disk inside the ISCO. If the black hole is rotating retrograde then the hole will become larger, and if the for prograde rotation the hole shrinks. The area of the hole will change by a factor of about 100 from smallest to largest. One method is called continuum fitting, and the other method is iron line method. In the second one one fits the iron line emitted from the disk.

What sets the jet power? We have magnetic flux $\Phi \sim Br_g^2$, and the gravitational radius $r_g = GM/c^2$. The jet power is approximated as

$$P_j \sim \Phi^2 (a/r_g)^2 = k\Phi^2 \quad (2.1)$$

I fell asleep again... Some “put money into blackhole get more money back” thingy. MAD takes energy and angular momentum from the black hole and put them into the jet, it also slows down BHs to a halt. What is MAD?

Now we want to talk about can jets die. At some point in time of simulation we turn on efficient radiative cooling. The disk collapses and breaks up into thin filaments. However the large scale magnetic flux only changes by a factor of a few, but the mass accretion rate increases to compensate, and the net energy flow is about the same.

How strong is B in Blazars? We can measure B in the jet using the core structure. The core is where the jet becomes transparent to its own synchrotron radiation $\tau_\nu \sim 1$. At higher ν the core will shift inward depending on the field strength

$$B \propto (dr_{\text{core}})^{3/4} \quad (2.2)$$

So one could use this to measure B by looking at the size of the core at two different frequencies. By the way this is a nature paper.