

THE CONNECTION BETWEEN STAR FORMATION RATE AND DARK MATTER HALO MASS IN THE EPOCH OF REIONIZATION

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ABSTRACT

We present updated constraints between the star formation rate and dark matter halo mass at redshift $z \sim 6$. The observational basis for our work is the luminosity function data obtained with from CFHTLS, HST Legacy Survey and UKIDSS. The constrains are based on an abundance matching methodology using data from cosmological N-body simulations in order to take into account the impact of cosmic variance in our estimates.

Subject headings: galaxies: high-redshift — methods: numerical

1. INTRODUCTION

Cosmological simulations still playing an important role on astrophysics and cosmology. They have become the bridge between observations and theory, a laboratory where is possible to create universes with different parameters and new physics.

This is an important tool in the study of dark matter. Since its existence was purposed, DM has been not observed directly neither on telescopes or high energy particle colliders, there exists many other theories explaining the observed universe behavior modifying the gravity itself.

The most successful and accepted cosmological model is the Lambda-CDM. Dark matter is a significative component of the universe. (?)

Hierarchy Structure Evolution: Early formation of small structures merging on major structures after.

From simulations, Halo Mass Function as function of redshift (or time).

In this paper is established a relation between dark matter and galaxy luminosity functions at high redshift. This relation is a direct connection between the Lambda-CDM model and the farthest galaxy observations made with space and grounded telescopes.

Star formation rate as function of time. Peak at $z \sim 2$.

Main Objective: Reproduce the observed luminosity function at redshift $z = 5.9$ from a DMH catalog from simulations.

This paper is organized as follows: In section 2... in section 3...

1.1. Halo Mass Function (HMF)

1.2. Cosmic Variance

2. LINKING GALAXY LUMINOSITY FUNCTION (GLF) AND DARK MATTER HALO (DMH) MASS

2.1. Magnitude to Luminosity

Luminosity Functions (LF) are usually expressed in terms of magnitude instead luminosity. Luminosity is the energy emmited by a source in a given wavelengt range, is a physical quantity.

Magnitude is a classification inherited from ancient

Greeks, this quantifies the response of the first astrometric device: the human eye, this perception grows loarithmically with the retrieved radiation.

Luminosity of any object can be compared with Sun Luminosity ($L_{\lambda\odot}$) at any wavelengt. With the Sun Magnitude as reference ($M_{\lambda\odot}$), the absolute magnitud of the objet at an specific wavelengt is given by:

$$M_{\lambda} = M_{\lambda\odot} - 2.5 \log_{10} \left(\frac{L_{\lambda}}{L_{\lambda\odot}} \right)$$

The solar absolute magnitude in the U filter is $M_{U\odot} = 5.61$, and the solar luminosity in the same filter is $L_{U\odot} = 10^{18.48} \text{ergs s}^{-1} \text{Hz}^{-1}$ or $L_{U\odot} = 3.02 \times 10^{18} \text{ergs s}^{-1} \text{Hz}^{-1}$. The solar luminosity can be used as refference unit, in this fashion, the typical luminosity of a galaxy can be expresed in terms of $10^8 - 10^{11}$ times the sun luminosity.

The absolute magnitude of a galaxy equation in the U filter:

$$M_U = 51.82 - 2.5 \log_{10}(L_U)$$

3. OBSERVATIONAL BACKGROUND

3.1. Star Formation Rate

The age of a star can be estimated by analyzing its spectrum. But when far galaxies are studied, individual stars can not be resolved. Is not possible to make a detailed census of the galaxy population. Only is possible to get information from the whole stellar population, an integrated spectrum.

There is a method(?) in wich a linear relation between SFR and luminosity in specific wavelength ranges can be assumed. This model allows to estimate the young stars fraction and the mean SFR over periods of $10^8 - 10^9 \text{yr}$ (?). The luminosity in the model, comes from the UV and the FIR broadband, also from specific recombination lines.

In a typical galaxy spectrum the visible wavelengths are dominated by the main sequence stars (A to early F) and G-K giants. In few wavelength ranges we have a significative contribution from the young stars rather than the old stars. The infrared and far infrared wavelengths emission is dominated by dust, this dust is heated by the whole stellar population, in particular by young, UV-bright stars (?).

On active galaxies the UV broadband emission is dominated by late-O and early-B type stars, with temperature

near to 40,000K. These hot and massive stars has a life-time below 10^7 Gyr, they spend their nuclear fuel faster than smaller and cooler sunlike stars.

The relation between UV luminosity and Star Formation Rate (??) is given by:

$$\text{SFR} (M_{\odot}\text{yr}^{-1}) = 1.4 \times 10^{-28} L_{\nu} (\text{erg s}^{-1}\text{Hz}^{-1}) \quad (1)$$

With Initial Mass Function (IMF) between $0.1M_{\odot}$ and $100M_{\odot}$, in the range of 1250 – 2500Å

The star forming rate will be:

$$\text{SFR} = k \times L_0 M \left[\left(\frac{M}{M_0} \right)^{-\beta} + \left(\frac{M}{M_0} \right)^{\gamma} \right]^{-1} \quad (2)$$

3.2. Dust Absorption

This paper is based on four main observational data sets from the Hubble Space Telescope and three ground-based telescopes. All $z \sim 6$ LBG candidates where discovered using the drop-out technique(?). All magnitudes are in AB system.

The data from the Hubble Space Telescope Legacy (HSTL) (?) is a compilation of observations since the installation of the Advanced Camera for Surveys (ACS) in 2002, through the near-infrared Wide Field Camera 3 (WFC3/IR) installed in 2009, up to 2012. The HST fields of view are: XDF, HUDF09-1, HUDF09-2, CANDELS-S/Deep, CANDELS-S/Wide, ERS, CANDELS-N/Deep, CANDELS-N/Wide, CANDELS-UDS, CANDELS-COSMOS and CANDELS-EGS, with areas of 4.7, 4.7, 4.7, 64.5, 34.2, 40.5, 62.9, 60.9, 151.2, 151.9 and 150.7 arcmin² respectively. The total area at redshift 6 corresponds to 740.8 arcmin² over five different lines of sight, with a total volume of $1.8 \times 10^6 \text{Mpc}^3$. Two cameras performed the observations: ACS and WFC3/IR, using B_{435} , V_{606} , i_{814} , z_{850} , I_{814} , Y_{098} , Y_{105} , J_{125} , JH_{140} and H_{160} filters. The limit magnitude is between $\sim 27.5\text{mag}$ in CANDELS-EGS and $\sim 30\text{mag}$ in the deepest field (XDF). Total number of $z=6$ LBG candidates is 940, most of them in the faint end of the LF, magnitudes in the rest frame are in the range $-22.52 \leq M_{1600} \leq -16.77$? calculated LF using a stepwise maximum-likelihood (SWML) based on ?. The Schechter parameters derived are: $\phi^* = (0.33^{+0.15}_{-0.10}) \times 10^{-3} \text{Mpc}^{-3}$, $M_{1600}^* = -21.16 \pm 0.20$ and $\alpha = -1.91 \pm 0.09$. ? reported that using just few fields of view, UVLF has a slightly non-Schechter-like form.

? worked also with HST, using the HUDF, CANDELS and GOODS fields, along with two of the Hubble Frontier Fields (HFF): deep parallel observations (unlensed fields) near the Abell 2744 and MACS J0416.1-2403 clusters. The HFF uses the ACS and the WFC3/IR with the same filters aforementioned but z_{850} . Total survey area is around $\sim 300\text{arcmin}^2$, with 706 photometric candidates at redshift 6 defined as the interval $5.5 < z < 6.5$. Total volume of this study is around $8 \times 10^5 \text{Mpc}^3$. The Schechter function parameters they found are $\phi^* = (1.86^{+0.94}_{-0.80}) \times 10^{-4} \text{Mpc}^{-3}$, $M_{1600}^* = -21.1^{+0.25}_{-0.31}$ and $\alpha = -2.02^{+0.10}_{-0.10}$.

? presented the sixth release of the Canada-France-Hawaii Telescope Legacy Survey CFHTLS. The observations where performed over four separated fields covering

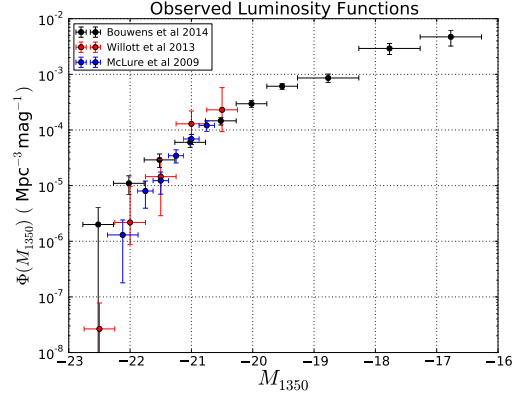


FIG. 1.— Observational data from ?? and ?.

a total area $\sim 4 \text{deg}^2$ (a large area), it gives this survey great robustness. Optical observations used MegaCam with $u^*g'r'i'z'$ filters. The main selection criteria: all the objects must be brighter than magnitude $z' = 25.3$. The final number of LBGs founded was 40. Moreover, they get spectroscopic confirmation for 7 candidates using GMOS spectrograph on the Gemini Telescopes, which has a $\ll 5.5$ -square arcmin field of view, giving a volume $\sim 1.09 \times 10^7 \text{Mpc}^3$. They show incompleteness in the sample due to foreground contamination and the detection algorithm; there is no warranty to have every object brighter than the limit magnitude on the faint limit. The full galaxy LF at $z = 6$ cannot be obtained as in other studies. Nevertheless, this survey was focussed on the highly luminous LBGs. LF is calculated using the stepwise maximum likelihood method of ?, within magnitudes from $M_{1350} = -22.5$ up to -20.5 . The luminosity function of $z = 5.9$ shows an exponential decline at the bright end, where feedback processes and inefficient last cooling limites star forming in bright galaxies hosted in the most massive halos.

? build the luminosity function for $z = 5$ and $z = 6$ using data from two ground-based telescopes: the United Kingdom Infrared Telescope in the near-IR imaging and Subaru Telescope for the optical imaging. They use the first data release of the UKIRT Infrared DeepSky Survey Ultra Deep Survey (UDS), together with the Subaru XMM-Newton Survey (SXDS). Total observed area is 0.63deg^2 uniformly covered by booth catalogues. The volume in this survey is $\sim 3 \times 10^6 \text{Mpc}^3$. The UKIRT was equipped with the WFCAM using JK filters. The Subaru was equipped with the Suprime-Cam with the $BVRi'z'$ filters. All candidates where brighter than $z' = 26$. The UV rest frame magnitude range is $-22.4 \leq M_{1500} \leq -20.6$. The LF was calculated using the maximum likelihood estimator of ?. Their analysis gave a total number of 104 LBG candidates in the redshift range $5.7 \leq z \leq 6.3$. LF was parameterized according to the Schechter function with $\phi^* = (1.8 \pm 0.5) \times 10^{-3} \text{Mpc}^{-3}$, $M_{1500}^* = -20.04 \pm 0.12$ and $\alpha = -1.71 \pm 0.11$.

The dataset was retrieved from ? graph using GAVO-DEXTER¹.

4. ABUNDANCE MATCHING METHODOLOGY

¹ <http://dc.zah.uni-heidelberg.de/dexter/ui/ui/custom>

The key element to connect SFR with halo mass from simulations is the observed UVLF. If there exist a function whom gives to each halo a UV luminosity then is posible to reproduce UVLF from DMH catalogs. The fitting parameters can be explored using a Markov Chain Monte-Carlo (MCMC) implementation. Once the parameters are found, the SFR-Halo Mass relation have been found.

We use the dark matter halo catalog from the Big MultiDark Plank 1 (MDPL) Simulation (?) with 2013 Planck cosmology (?). MDPL is quite similar to the Big Bolshoi (1Gpc h⁻¹) (?) and its predecesor Bolshoi (?) (250Mpc h⁻¹), booth of them with WMAP5 cosmology, but MDPL has bigger mass resolution. Those halo catalogs are available at the MultiDark Database² (?).

The MDPL run is a N-body dark matter only simulation based on the L-Gadget2 code. The simulated volume is a cubic box of 1Gpc h⁻¹ side length. It has 3840³ dark matter particles of 1.51 × 9M_⊙h⁻¹ mass each one. The 2013 Planck cosmology is defined by the following parameters: Ω_M = 0.307, Ω_B = 0.048, Ω_Λ = 0.730, σ₈ = 0.829, n_s = 0.96 and H₀ = 67.8. The DMH Catalog at z = 6 contains ~ 10.9 × 10⁷ halos, to avoid incompleteness in the low mass end, halos with mass below 10^{10.3}M_⊙h⁻¹ are rejected. In order to study how cosmic variance can affect measurements, the original catalog is divided into 64 small cubic boxes, with a similar volume to the observations.

The following treatment is applied to every one small box.

The halo abundance matching technique has been widely used (??) resulting in good reproduction of the observed galaxy clustering. (??). Here we use the simplest case with few premises:

1. Each halo in the catalog hosts one galaxy. There are not empty halos, also none of halos has two or more galaxies.
2. The UV luminosity of each galaxy is function of one variable; the mass of the DMH in which is located. This function must be monotone, this guarantees that most massive halos will host the most luminous galaxies in the same volume.

4.1. Fitting Model

The observed UV Luminosity Function has two slopes. To reproduce it we decide to use the four parameter model:

$$L_{UV}(M) = L_0 M \left[\left(\frac{M}{M_0} \right)^{-\beta} + \left(\frac{M}{M_0} \right)^{\gamma} \right]^{-1} \quad (3)$$

where M is the hosting DMH mass, L_0 is a normalization constant, M_0 is the critical mass where the luminosity function has a slope change, β and γ are the slopes. This equation has a similar fashion to the mass to light relation (?) and the mean relation between stellar mas of a galaxy and the mass of its halo used by ?.

4.2. Dust Attenuation

² <http://www.multidark.org>

Dust in star forming galaxies absorbs part of the UV radiation and reemits on IR. The more massive is the galaxy then more dust contains and dust attenuation will be greater. The relation between dust attenuation and magnitude have been already studied, with this relation we can infer the dust-free UV luminosity of a galaxy from observations, resulting in a more accurate inferred SFR.

The UV Spectral Slope β was introduced by ? as a UV color to study dust attenuation in local starburst galaxies and extrapolating to high redshift galaxies. This index appears when a power law fitting is performed over the spectral flux f as function of the wavelength λ ;

$$f \propto \lambda^{\beta}.$$

The relation for ultra-violet attenuation at 1600Å they found is:

$$A_{1600} = 4.43 + 1.99\beta, \quad (4)$$

with A_{1600} in magnitude units.

Due LBGs have more similar spectra properties to local starburst galaxies rather than AGNs for example, we can assume that local calibration of $\beta - A_{1600}$ can be applied also for LBGs, “The main requirement is that the data include fluxes in two broad bands or coarse spectra covering the rest-frame UV.”

? uses the fluxes on different bands to estimate β on each LBG candidate found with $z \sim 4 - 7$. After in redshift groups they found a linear relation between β and the UV magnitude:

$$\langle \beta \rangle = \frac{d\beta}{dM_{UV}} (M_{UV,AB} + 19.5) + \beta_{M_{UV}=-19.5} \quad (5)$$

with $\beta_{M_{UV}=-19.5} = -2.20$ and $d\beta/dM_{UV} = -0.21$ at $z = 5.9$.

? used the aforementioned relations to infer the corrected luminosity functions, i.e. dust-free luminosity functions, and the corrected SFR at $z = 4$.

Here we use the inverse relation, starting from the intrinsic or dust-free galaxy magnitude, obtaining the observed magnitude:

$$M_{obs} = \begin{cases} \frac{M_{int}-4.616}{1.259}, & \text{if } A > 0 \\ M_{int}, & \text{else} \end{cases} \quad (6)$$

4.3. MCMC

We used a Markov Chain Monte Carlo method to find the best parameters and its uncertainties over each one of the boxes and each observational dataset. The code was written on Python using the Scipy, Numpy and Matplotlib libraries.

First of all, one of the four observational data sets (Bouwens, Finkelstein, Willott and McLure) is selected to be the model to fit. From the whole simulated halo catalog, a subsample with cubic box shape of 250Mpc h⁻¹ length is selected.

With an initial set of parameters (α, γ, M_0 , and L_0) the UV luminosity for the halo is calculated as function of his mass according to equation 3. Then Luminosity is converted to magnitude units using:

$$M_{UV} = 51.82 - 2.5 \log_{10}(L_{UV}).$$

If we consider the dust attenuation im the equation 6 we have a dust-corrected UV Luminosity Function.

The luminosity function is constructed as an histogram of the magnitudes normalized by the volume of the catalog. Each observed luminosity function has a different bin range.

Once having the LF, we compare our LF against observed LF. The error function we consider is the sum of the square difference over each bin, divided by the observational data uncertain.

$$\chi^2 = \sum_{i=0}^n \frac{(x_{i,obs} - x_{i,fit})^2}{2\sigma_i^2}$$

We worked on the logarithmic space of luminosities to have a good fitting on six decades.

The likelihood will have this property:

$$\mathcal{L} \propto \exp\left(-\frac{\chi^2}{n}\right)$$

, where n is the number of degrees of freedom. We have the maximum likelihood when the error is the minimum.

Each MCMC step will have a small variation of the parameters, the UV luminosity and magnitude are calculated again for each halo, we have a new error χ_{new}^2 and a new likelihood \mathcal{L}_{new} .

Following the Metropolis method, we compare the new and the old likelihood:

$$R = \frac{\mathcal{L}_{new}}{\mathcal{L}_{old}} = \exp(\chi_{new}^2 - \chi_{old}^2)$$

If $R \geq 1$, then we immediately accept the new set of parameters and start the next MCMC step.

Else, we have a chance to keep the new set. When $R < 1$, we compare with a uniformly random number p in the range $[0, 1]$. if $R > p$ we accept the new set of parameters and start the next MCMC step. Else we reject the new set and start again with the old set.

We performed 10.000 effective MCMC steps, plus 1.000 thermalization steps over each box. We repeated for the same box without consider dust attenuation.

Then we perform the same method over the 64 boxes and the resting three data sets.

The restrictions imposed over the parameters where $0 \leq \alpha \leq 2.0$ and $\gamma \geq 0$.

Finally, the UV luminosity for each galaxy can be directly related with SFR according to ?. This model is accurate within the range of $10^8 - 10^9 \text{yr}^{-1}$ (?). The relation between UV luminosity and Star Formation Rate (??) is given by:

$$\text{SFR} (M_{\odot} \text{yr}^{-1}) = 1.4 \times 10^{-28} L_{\nu} (\text{erg s}^{-1} \text{Hz}^{-1}) \quad (7)$$

4.4. The Luminosity Model

In this model we have made two assumptions:

1. Each halo in the catalog hosts one galaxy. There are not empty halos, also none of halos has two or more galaxies.
2. The UV luminosity of each galaxy is function of one variable: the mass of the DMH in which is located.

The simplest relation we can have is a powerlaw:

$$L = L_0 M^{\alpha} \quad (8)$$

but has not well agreement with observed luminosity functions.

A better model is a four parameter function. Each galaxy has a luminosity given by:

$$L = L_0 M \left[\left(\frac{M}{M_0} \right)^{-\beta} + \left(\frac{M}{M_0} \right)^{\gamma} \right]^{-1} \quad (9)$$

where M is the hosting DMH mass, L_0 is a normalization constant, M_0 is the critical mass where the luminosity function has a slope change, β and γ are the slopes. This equation has a similar fashion to the mass to light relation (?) and the mean relation between stellar mass of a galaxy and the mass of its halo used by ?.

There are more complex models(?) that includes a random behavior: galaxies has not synchronization on the beginning of star forming stage, also this stage may be time limited. This is called duty cycle. It is probable to have in the observations some invisible galaxies in the UV continuum due their duty cycle may has not started as well it may ended. Also may be present a normal distribution of the luminosity around the expected values.

5. RESULTS

Best fit parameters:

Cosmic Variance Dust Attenuation model:

$\log_{10}(L_0)$, mean value = 18.0668181805, cosmic std dev=0.0288558407431, mean sigma per box =0.115498563539

$\log_{10}(M_0)$, mean value = 11.1894829902, cosmic std dev=0.0718242117247, mean sigma per box =0.72909362196

β , mean value = 1.48014744519, cosmic std dev=0.175026973523, mean sigma per box =1.2988714515

γ , mean value = 0.402845130968, cosmic std dev=0.0575764062455, mean sigma per box =0.314709555486

Cosmic Variance No-Dust Attenuation model:

$\log_{10}(L_0)$, mean value = 17.8076246968, cosmic std dev=0.0198125186745, mean sigma per box =0.150082796188

$\log_{10}(M_0)$, mean value = 11.119460579, smic std dev=0.0601358065131, mean sigma per box =0.642532804052

β , mean value = 1.43705896343, cosmic std dev=0.178253100134, mean sigma per box =1.37416358513

γ , mean value = 0.518242492031, cosmic std dev=0.0461244563387, mean sigma per box =0.245817572122

6. DISCUSSION

7. CONCLUSIONS

ACKNOWLEDGMENTS

Acknowledgments...

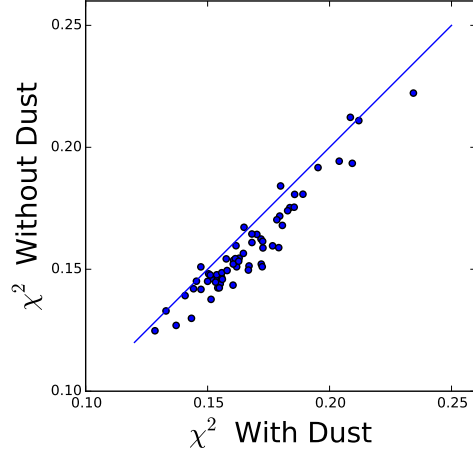


FIG. 2.— Best fit to Willott comparison between the two models. Each point represent χ^2 calculated over each small-box. The solid line represents the ratio 1:1

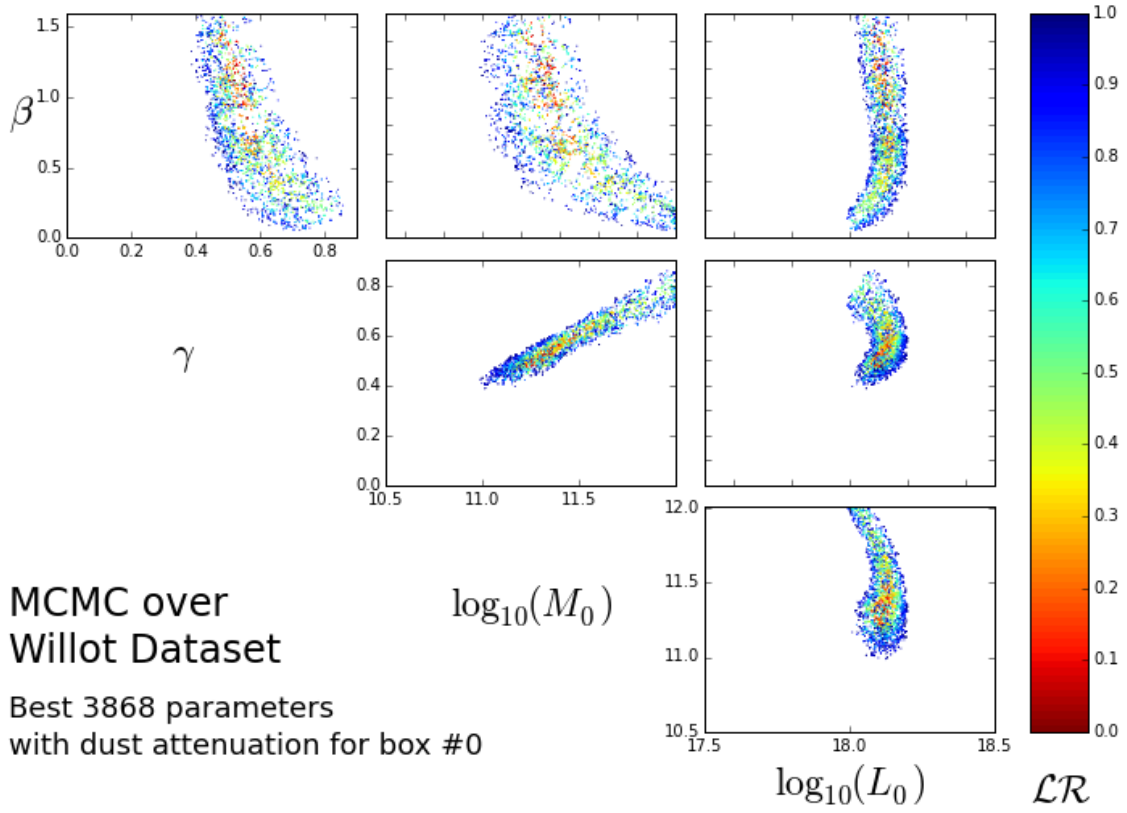


FIG. 3.— Parameter dispersion fitting the Willott with the Dust Attenuation model. 1σ is defined by the likelihood ratio between 0.0 (red) and 0.5 (green)

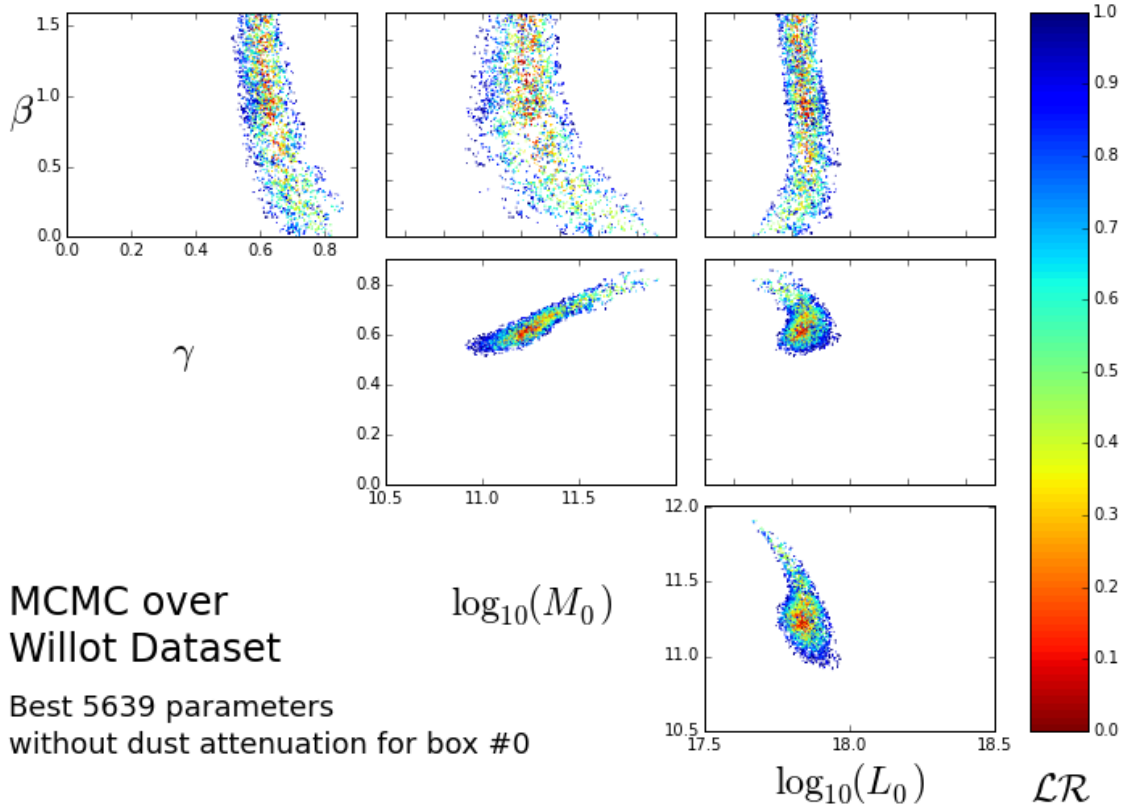


FIG. 4.— Parameter dispersion fitting the Willott with the No-Dust Attenuation model. 1σ is defined by the likelihood ratio between 0.0 (red) and 0.5 (green)

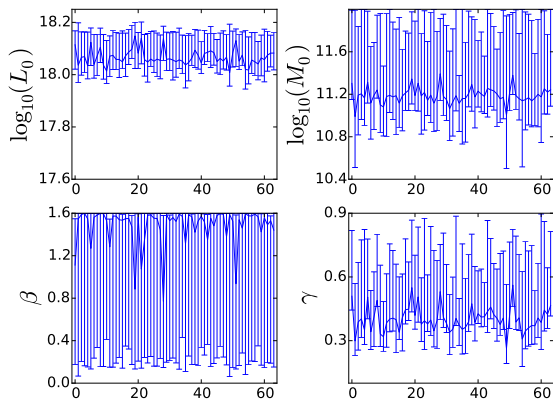


FIG. 5.— Individual small-box parameter estimation with dust attenuation: Best fit values to the Willott data set (solid line) and 1σ confidence interval using likelihood ratio $\mathcal{LR} = 0.5$. The x axis corresponds to the box number.

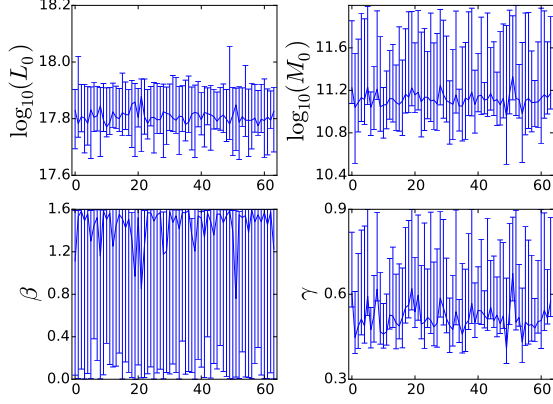


FIG. 6.— Individual small-box parameter estimation with no-dust attenuation: Best fit values to the Willott data set (solid line) and 1σ confidence interval using likelihood ratio $\mathcal{LR} = 0.5$. The x axis corresponds to the box number.

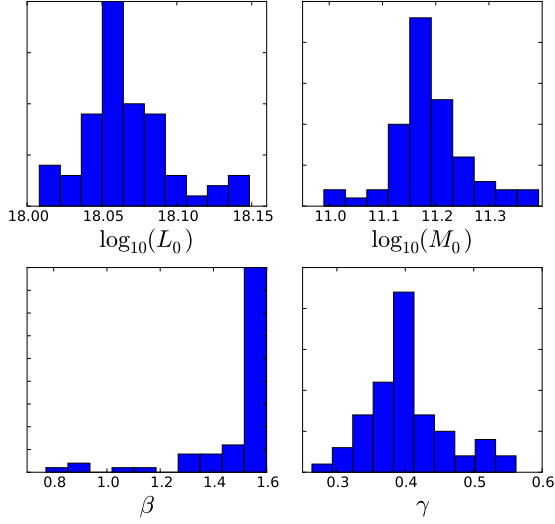


FIG. 7.— Best fit parameter distribution due to cosmic variance with Willott data in the dust attenuation model.

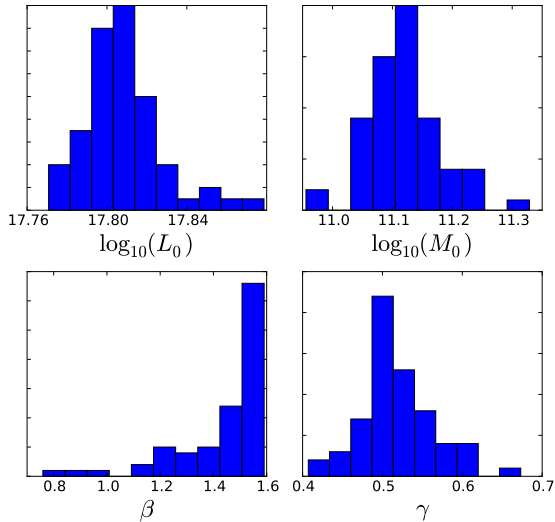


FIG. 8.— Best fit parameter distribution due to cosmic variance with Willott data in the no-dust attenuation model.

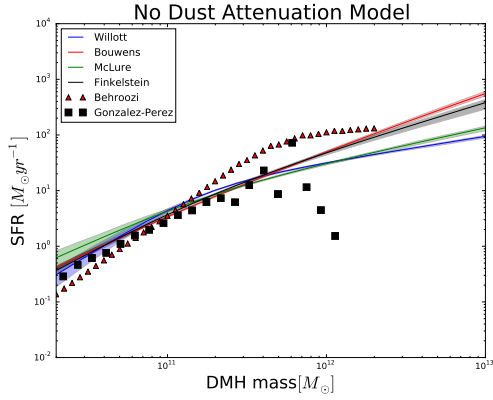


FIG. 9.— Star formation rate as function of the dark matter halo mass without dust attenuation. Solid lines represents the mean SFR value over the small boxes within 50% shaded region. Comparison with the GALFORM semi-analytical model (?) and a implementation of abundance matching model (?).

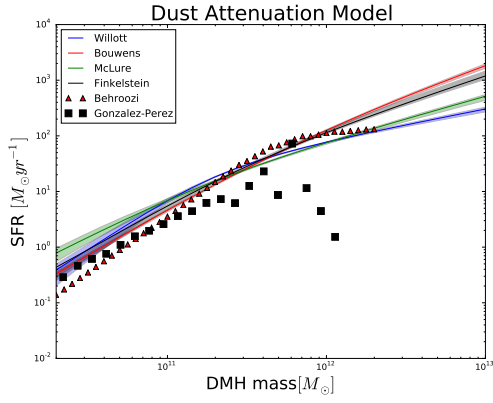


FIG. 10.— Star formation rate as function of the dark matter halo mass with dust attenuation. Solid lines represents the mean SFR value over the small boxes within 50% shaded region. Comparison with the GALFORM semi-analytical model (?) and a implementation of abundance matching model (?).