

THE PLACE OF THE LOCAL GROUP IN THE COSMIC WEB

J. E. FORERO-ROMERO¹ AND R. GONZÁLEZ²

¹ Departamento de Física, Universidad de los Andes, Cra. 1 No. 18A-10, Edificio Ip, Bogotá, Colombia

² Instituto de Astrofísica, Pontificia Universidad Católica, Av. Vicuña Mackenna 4860, Santiago, Chile

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ABSTRACT

We explore the characteristics of the Local Group (LG) location in the cosmic web using a cosmological simulation in the Λ CDM cosmology. We use the Hessian of the gravitational potential to define the cosmic web, where regions are classified as either peak, sheet, filament or void. The LG in the simulations are split into two samples. The first is a general sample composed by halo pairs with similar masses and isolation criteria as observed in the LG. The second is a subset with additional kinematic constraints. We find that the pairs in the LG sample with all constraints are: i) Preferentially located in filaments and sheets, ii) located in a narrow range of local overdensity $0 < \delta < 4$, web ellipticity $0.1 < e < 1.0$ and prolateness $-0.5 < p < 0.5$. iii) Clearly aligned with the cosmic web, in particular for pairs in filaments/sheets the orbital angular momentum tends to be perpendicular to the filament direction or the sheet plane. A stronger alignment is present for the vector linking the two halos, which lies along the filament or the sheet plane. We show that the first and second results are expected trends with the LG total mass. However, the strong alignment with the cosmic web cannot be explained in the same way.

Subject headings: galaxies: Local Group — dark matter

1. INTRODUCTION

The spatial and kinematic configuration of the Local Group (LG) galaxies is rare to find in the local Universe and in cosmological simulations. The LG is dominated by the two big spirals MW and M31, the next most-luminous galaxy is M33 which is ~ 10 times less massive than M31, followed by several less luminous dwarf galaxies, up to a distance of ~ 3 Mpc. The velocity vector of M31, with a low tangential velocity is consistent to a head-on collision orbit toward the MW (Cox & Loeb 2008; van der Marel et al. 2012b; Sohn et al. 2012).

Another feature of the Local group is the relatively low velocity dispersion of nearby galaxies up to ~ 8 Mpc (Sandage & Tammann 1975; Aragon-Calvo et al. 2011, and references therein). Environment around the Local Group has density quite close to the average density of the universe (Klypin et al. 2003; Karachentsev 2005). In addition, the closest massive galaxy cluster, the Virgo Cluster, is ≈ 16.5 Mpc away (Mei et al. 2007).

All this combination of features make LG analogues rare to find. Using numerical simulations González et al. (2013a) found less than 2% MW-sized halos reside in a pair similar to MW-M31 and in a similar environment. Furthermore, if we select pairs constrained within 2σ error from current observational measurements of the velocity components and distance to M31, there are only 46 systems in a cubic volume of $250 h^{-1}$ Mpc side, giving a number density $\sim 1.0 \times 10^{-6}$ Mpc³, comparable to the abundance of massive clusters.

Forero-Romero et al. (2013) also studied MW-M31 pairs in numerical simulations finding the typical quantities characterizing the orbital parameters of the LG are rare among typical pairs, but not enough to challenge the Λ CDM model.

Another definition of LG analogues is made by Li & White (2008), but despite differences occur in the definitions and resulting fraction of LG analogues, all are in agreement with a low frequency of these pairs.

To better understand the properties of the LG and how this uncommon pair configuration fit in the cosmological context, an immediate question arise. What else can we say of the LG at larger scales?. In particular, which are the typical/preferred locations of these systems within the Cosmic Web?. To what extent is this an expected configuration in Λ CDM.

Looking the LG at larger scales we have it is located in a diffuse and warped filament/wall connecting Virgo Cluster with Fornax Cluster, some nearby galaxies and groups members of this large structure are the Maffei group, NGC 6744, NGC 5128, M101, M81, NGC1023, Cen A group. At this scale, there is no evident alignment of the MW-M31 orbital plane with the local filament or the Virgo-Fornax direction. However, if we look in a smaller volume below scales of ~ 6 Mpc, there is a clear alignment of the MW-M31 orbit with a local plane as shown by Figure 3 in Courtois et al. (2013).

In this paper we study the large scale environment of LG analogues in the context of Λ CDM. We use the Bolshoi simulation to explore in what structures they reside and if there is any correlation or alignment with the cosmic web. The large scale Environment is defined by the cosmic web components identified by Forero-Romero et al. (2009), and we use the LG analogues computed by González et al. (2013a). We pay special attention to quantify the kind of environment that hosts LG pairs and their alignments with respect to the preferred directions defined by the T-web.

This paper is organized as follows. In Section 2 we present the N-body cosmological simulation and the algorithm to define the cosmic web, next in 3 we describe

the sample of LG analogues extracted from the simulation. In Section 4 we presents our results to continue with a discussion about their implications for our Local Group in Section 5 to finally conclude in Section 6.

2. SIMULATION AND WEB FINDING ALGORITHM

2.1. The Bolshoi simulation

We use the Bolshoi simulation of Λ CDM cosmology: $\Omega_m = 1 - \Omega_\Lambda = 0.27$, $H_0 = 70$ km/s/Mpc, $\sigma_8 = 0.82$, $n_s = 0.95$ (Klypin et al. 2011), compatible with the constraints from the WMAP satellite (Hinshaw et al. 2013). The simulation followed the evolution of dark matter in a $250h^{-1}$ Mpc box with spatial resolution of $\approx 1h^{-1}$ kpc and mass resolution of $m_p = 1.35 \times 10^8 M_\odot$. Halos are identified with the BDM algorithm (Klypin & Holtzman 1997). The BDM algorithm is a spherical overdensity halo finding algorithm and is designed to identify both host halos and subhalos.

2.2. Cosmic web identification

The web finding algorithm is based on the tidal tensor computed as the Hessian of the gravitational potential field

$$T_{ij} = \frac{\partial^2 \phi}{\partial r_i \partial r_j}, \quad (1)$$

where r_i , $i = 1, 2, 3$ refers to the three spatial comoving coordinates and ϕ is the gravitational potential renormalized to follow the Poisson equation $\nabla^2 \phi = \delta$ where δ is the matter overdensity.

This tensor is real and symmetric, which means that can be diagonalized. We note its eigenvalues as $\lambda_1 \geq \lambda_2 \geq \lambda_3$ and their corresponding eigenvectors \hat{e}_1 , \hat{e}_2 and \hat{e}_3 . The web classification compares each one of the three eigenvalues to a threshold value λ_{th} . If the three, two, one or zero eigenvalues are larger than this threshold the region is classified as peak, filament, sheet or void, respectively.

Forero-Romero et al. (2009) performed a detailed study for the topology of the cosmic web and its visual counterpart as a function of the parameter λ_{th} . They found that reasonable results in terms of the volume fraction occupied by voids, the visual inspection and the halo populations in each web type can be reached by values of $0.2 < \lambda_{th} < 0.4$. In this paper we choose the value of $\lambda_{th} = 0.3$ to proceed with our analysis. This is only relevant to the classification of the simulation into web elements. Other results are completely independent of this choice. Nevertheless we have checked that the main conclusions of this work do not depend on the choice of λ_{th} .

The algorithm to compute the potential is grid based. First we interpolate the density into a cubic grid with a Cloud-In-Cell (CIC) scheme and smooth it with a gaussian kernel. Then we obtain the gravitational potential using FFT methods and use finite differences to compute the Hessian at every point in the grid. In our case we have used a grid size on and a gaussian smoothing with two times larger as the typical separation between the two halos in the Local Group. The purpose of this choice is to have both halos in the pair a common environment. In this paper we use a grid spacing of $s = 0.97$

h^{-1} Mpc, corresponding to a 256^3 grid in the Bolshoi volume. The scale for the gaussian smoothing uses the same value.

We use the matter overdensity, ellipticity and the prolateness to further characterize the web. These quantities are defined in terms of the eigenvalues as follows

$$\delta = \lambda_1 + \lambda_2 + \lambda_3, \quad (2)$$

$$e = \frac{\lambda_3 - \lambda_1}{2(\lambda_1 + \lambda_2 + \lambda_3)}, \quad (3)$$

$$p = \frac{\lambda_1 + \lambda_3 - 2\lambda_2}{2(\lambda_1 + \lambda_2 + \lambda_3)}. \quad (4)$$

We also measure the alignment of the LG halos with respect to the cosmic web defined by the eigenvector. To this end we characterize each LG pair by two vectors. The first is \hat{n} , the vector marking the axis along the orbital angular momentum of the pair, normal to its orbital plane; the second is \hat{r} , the vector that connects the halos in the pair. We quantify the alignment using the absolute value of the cosinus of the angle between the two vectors of interest $\mu = |\hat{e}_i \cdot \hat{n}|$ or $\mu = |\hat{e}_i \cdot \hat{r}|$, where $i = 1, 2, 3$.

The data of the BDM halos and the Tweb is publicly available through a database located at <http://www.cosmosim.org/>. A detailed description of the database structure was presented by Riebe et al. (2013).

3. LOCAL GROUP ANALOGUES

To construct a sample of the MW-M31 pairs at $z \approx 0$, we use a series of simulation snapshots at $z < 0.1$ (since the last ≈ 1.3 Gyr) spaced by $\approx 150 - 250$ Myr. This is done because a particular configuration of MW and M31 is transient and corresponds to a relatively small number of systems at one snapshot. By using multiple snapshots we can increase the sample of systems in such configuration during a period of time in which secular cosmological evolution is small.

The LG analogues or General Sample (GS) in this paper are pairs selected in relative isolation, and in a wide range of masses from $M_{200c} = 5 \times 10^{11} M_\odot$ to $5 \times 10^{13} M_\odot$. Isolation criteria includes a pair closer than 1.3 Mpc, and with no massive neighbors within 5 Mpc. In addition we require that pairs have no Virgo-like neighbor halo with mass $M_{200c} > 1.5 \times 10^{14} M_\odot$ within 12 Mpc. We have 5480 pairs under these general criteria. A full description of the selection criteria can be found in González et al. (2013a,b).

We also define two subsamples more closely related to the MW-M31 dynamics according to the tolerance in additional constraints. A sample named 2σ , correspondig to LG analogues constrained by two times the observational errors in the orbital values (radial velocity, tangential velocity, and separation), and a more relaxed sample named 3σ for LG analogues constrained by three times observational errors accordingly. The number of pairs in each sample is 46 and 120 respectively, notice we have less pairs than in González et al. (2013a) results, since we removed pairs which are too close at $z = 0$, i.e. their virial radii overlaps, also we removed a couple pairs that merged or change their mass more than 20% at present time since they were detected at $z < 0.1$.

Sample	Peak n (%)	Filament n (%)	Sheet n (%)	Void n (%)
2σ	4 (8.7)	24 (52.2)	17 (36.7)	1 (2.2)
3σ	10 (8.3)	58 (48.3)	47 (39.2)	5 (4.2)
General	1312 (23.9)	1472 (26.9)	1769 (32.3)	927 (16.9)
General ($12.1 < \log_{10} M_{\text{LG}}/M_{\odot} < 12.3$)	8 (1.4)	334 (55.5)	259 (43.0)	1 (0.1)

TABLE 1

NUMBER OF PAIRS IN THE FOUR DIFFERENT KINDS OF ENVIRONMENTS FOR EACH OF THE THREE SAMPLES PRESENTED IN SECTION 3. IN PARENTHESIS THE SAME NUMBER AS A PERCENTAGE OF THE TOTAL POPULATION. THE LAST LINE IN THE TABLE CORRESPONDS TO THE GENERAL SAMPLE WITH AN ADDITIONAL MASS CUT FOR THE TOTAL PAIR MASS.

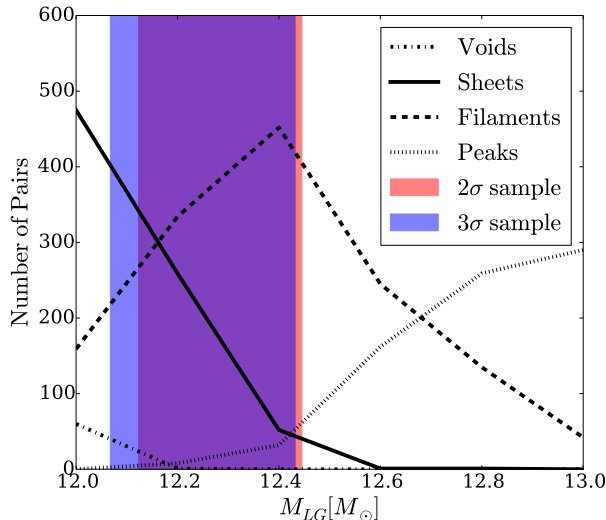


FIG. 1.— Mass distribution of pairs in the different environments for the general sample. The shaded regions show the mass ranges of 2σ and 3σ samples.

4. RESULTS

4.1. The preferred environment for LGs

The first result we explore is the kind of environment occupied by our LGs. We find that the LGs in the general sample are located across all different environment without any strong preferences; 1/3 are located in sheets, 1/4 in peaks, 1/4 in filaments and the remaining 1/6 in voids.

The situation in the restricted 2σ and 3σ samples is very different. By large the LGs in these samples are located in filaments and sheets. In both samples, $\sim 50\%$ of the pairs can be found in filaments while $\sim 40\%$ are in sheets. These absolute numbers in each environment for each sample are presented in Table 1.

This difference between the general and the restricted samples is due the mass ranges covered by each sample. In González et al. (2013a) the mass range covered by 2σ and 3σ is very narrow and it is used to constraint the LG mass. We show in table 1 that subset of the GS having a similar mass range to 2σ and 3σ reproduces similar environment fractions.

Figure 1 clearly shows the correlation between environment and total pair mass. Each line represents the mass distribution of pairs in the four different environments for the general sample. High mass pairs tend to be located in peaks and filaments while less massive ones in voids and sheets. The shaded regions represent the 68%

confidence intervals of the mass distributions of 2σ and 3σ samples.

4.2. Web Overdensity, Ellipticity and Anisotropy

We now describe the preferred place of the LG samples in terms of the web overdensity, ellipticity and anisotropy as defined in Section 2.

Figure 2 shows dependency of the web overdensity, ellipticity, and prolateness on pair mass for the different samples. GS is represented by the solid lines with the associated errors covered by the shaded region. The symbols represent the results for the 2σ and 3σ samples. In all cases it is immediately clear that the range of values for the 2σ and 3σ samples are completely expected from its mass dependence.

Left panel shows the overdensity dependence on pair mass. Higher mass pairs are located in high density regions. The 2σ and 3σ samples having a narrower mass range as shown in previous figure, are consequently located within a narrower range of overdensities $0.0 < \delta < 4.0$ peaking at $\delta \sim 1$. This is also consistent with the fact that these samples are mostly found in filaments and sheets. The average overdensity of 2σ and 3σ samples is expected from the values in the GS within the same mass range.

Middle and right panels show web ellipticity and absolute prolateness dependence on mass. Again we noticed that within the same mass range, the 2σ and 3σ average ellipticity and prolateness does not differ significantly from GS. For the 2σ and 3σ samples most of the pairs are located in a narrow range for ellipticities in the range $0.1 < e < 1.0$, and prolateness $0.5 < p < 0.5$.

4.3. Alignment along the cosmic web

We explore the alignment of the pairs with respect to the cosmic web. We define the unit vector \hat{n} , normal to the pair orbital plane computed from the pair positions and relative velocity. Then, we compute the dot product of \hat{n} and each one of the three eigenvectors for the three samples.

In Figure ?? we show the main result of that study, it presents the cumulative distribution of $\mu \equiv \hat{e}_i \cdot \hat{n}$ for the three eigenvectors $i = 1, 2, 3$. Lines indicate the different samples. The straight line corresponds to the expected result for a random vector distribution.

From this Figure we can see two important features: First, there is a strong anti-alignment signature between \hat{n} and the third eigenvector, and this signature reverses with the first eigenvector. However, in the case of the second eigenvector, no significant deviations are found. Second, for \hat{e}_3 vector the anti-alignment become stronger for the samples more closely related to the LG, then for

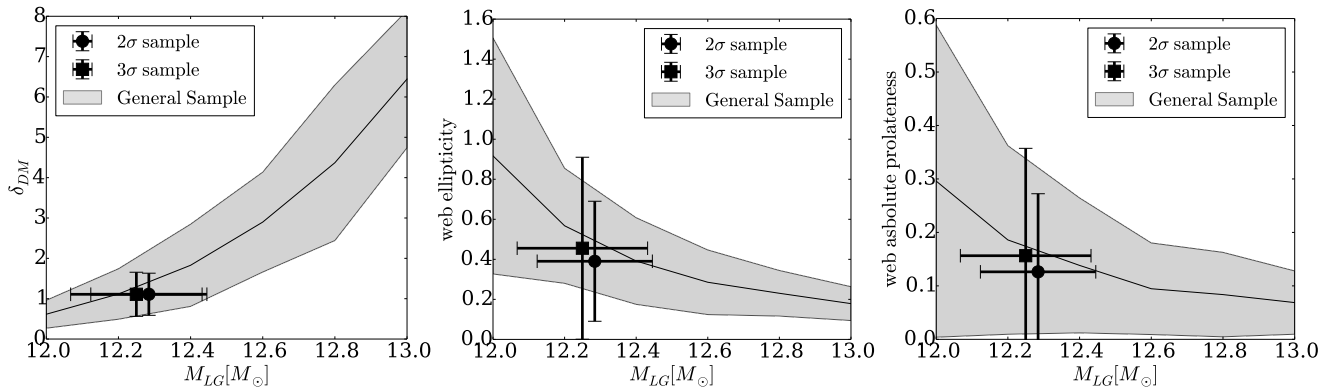


FIG. 2.— Mass dependency of the average dark matter overdensity (left), web ellipticity (middle) and web absolute value prolateness (right) at the pair location. 2 σ it is stronger than for 3 σ , and the latter is stronger than for General sample.

Quantitatively, the anti-alignment feature found with the \hat{e}_3 vector means that for 2 σ sample, $\sim 50\%$ pairs have $\mu < 0.2$, and $\sim 75\%$ pairs have $\mu < 0.4$. We also test that does not change significantly on environment type, in particular these trends hold very well for pairs in filaments and walls.

If we consider only pairs in filaments, we have that the pair orbit tends to be perpendicular to the filament direction, and both pair members lie along the filament, if we recall the 3 σ -2 σ samples are constrained with a mostly radial velocity component, then the pairs are moving into the filament. In the case of sheets, we have the pair orbital plane tend to lie within the sheet plane.

These alignment features are in agreement with the scenario that pairs created in-situ or falling into a filament/wall align their orbits with the large scale structure in a relaxation process where pair members tend to moves along the slowest collapsing directions. However, the most intriguing feature is the dependence of this alignment with how close we represent the LG, is still not clear. We explore if the pair masses are responsible of this, but we found no dependence at all.

In Figure 5 we show the relation between the median of $\mu \equiv \hat{e}_i \cdot \hat{n}$ and mass for the general sample, and found no significant relation. Lines in the figure show the median μ -mass relation for the three eigenvectors in the General sample. Dots represent the median μ for the three eigenvectors in the 2 σ -3 σ samples. In the case of $\|\hat{e}_3 \cdot \hat{n}\|$, the figure clearly shows for a fixed mass range around $1.5 - 2.0 \times 10^{10} M_\odot$, that the median μ decreases as we better constraint the LG samples, being larger for General sample and smaller for 2 σ sample accordingly.

4.4. Spin and Orbital Angular Momentum

In 4.2 we defined the unit vector \hat{n} , normal to the pair orbital plane which can be asociated with the orbital angular momentum direction \hat{J} . There, we show several alignment features of the pair orbital momentum with the environment. In this subsection we explore the alignment of the individual spin angular momentum of each pair member \hat{j}_A and \hat{j}_B with each other and with the orbital angular momentum.

We show in figure 6 the cumulative distribution of dot

product between the spin of both members in the pair. We found a slight alignment of spin vectors for the 2 σ sample. However, we found no significant alignment between spin and orbital angular momentum for any sample.

In the LG, the angle between MW and M31 spin is $\sim 60^\circ$ and the angles between spins and orbital angular momentum are $\sim 33^\circ$ and $\sim 76^\circ$ for MW and M31 respectively (van der Marel et al. 2012a). Then the LG show no particular alignment aswell.

5. DISCUSSION

We explored the characteristics of the LG location in the cosmic web.

LG pairs are preferentially located in filaments and sheets. The mass range of the pairs sample plays an important role in the fractions for each environment type as shown in table 1 and figure 1.

There is a clear anti-alignment between the third eigenvector and the vector normal to the pair orbital plane. This feature is not mass dependant as can be shown in figure 3. For pairs in filaments this means the pair orbit tends to be perpendicular to the filament direction, and both pair members lie along the filament, and in the case of pairs in sheets, we have the pair orbital plane lie in the sheet plane. We also found this anti-alignment becomes stronger for the samples being more closely related to the LG, it means it is stronger for 2 σ sample. All these alignment features are consistent with pair orbits relaxing their movements towards the slowest collapsing directions at larger scales.

LG pairs are located in a narrow range of local overdensity, ellipticity and prolateness. And this is consistent with the narrow mass range selection of specific samples.

The LG spin and orbital angular momentum have no particular alignment, and we see no significant alignments in our LG pairs, however for the alignment between pair spins we found a slight alignment signature for 2 σ sample as can be seen in figure 5.

6. CONCLUSIONS

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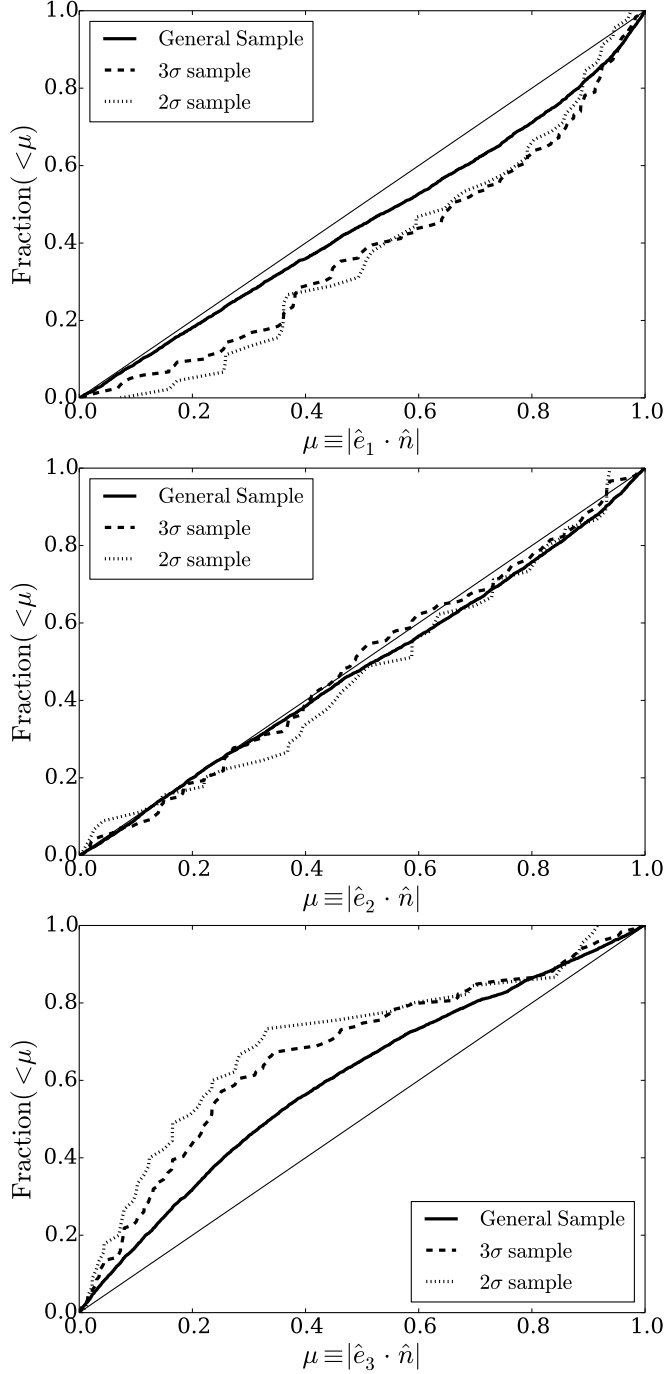


FIG. 3.— Cumulative distributions for the alignment between the normal vector to the pair orbital plane, \hat{n} , and the three eigenvectors in the Tweb.

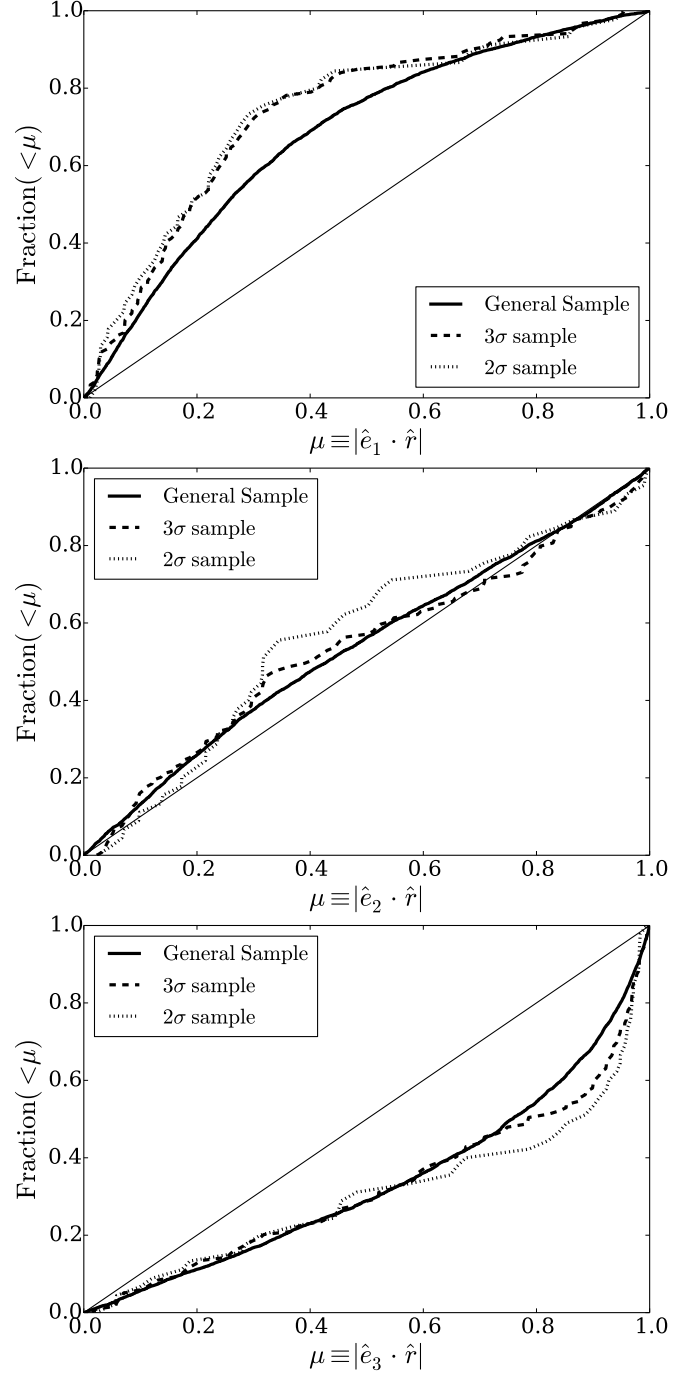


FIG. 4.— Cumulative distributions for the alignment between the vector linking the two halos in the pair, \hat{r} , and the three eigenvectors in the Tweb.

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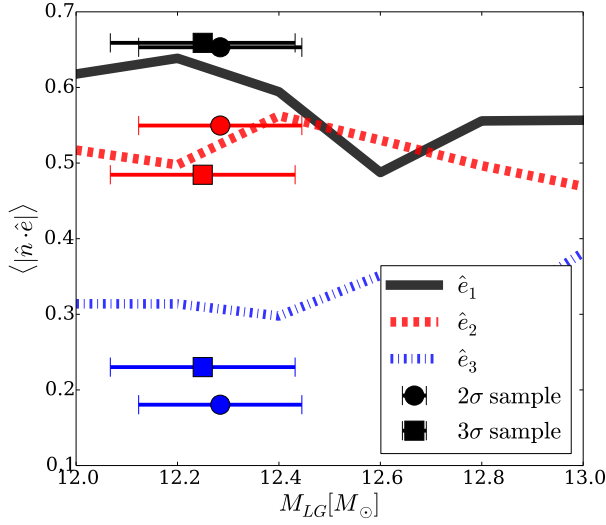


FIG. 5.— Mass dependency of the median value for the dot product between the normal vector \hat{n} and each one of the eigenvectors. The lines show the trends for the general sample.

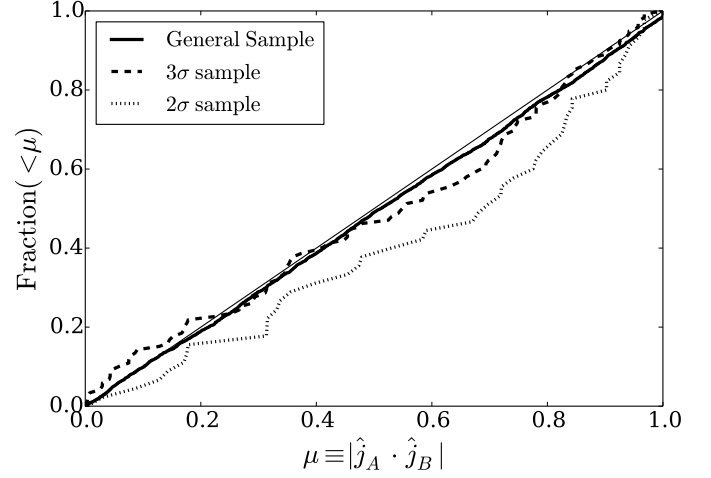


FIG. 6.— Alignment between the two angular momentum vectors of the two halos in the pair.

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