

INTRODUCTION

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ABSTRACT

Galaxies must accrete gas from the intergalactic medium (IGM) in order to sustain star formation at observed levels. In order to understand this complex process, and how it influences galaxy evolution globally, it is necessary to understand the physical conditions and distribution of the gas around galaxies, known as the circumgalactic medium (CGM). This thesis aims to address these questions through the largest-to-date survey of low column density Ly α absorption detected in the spectra of background QSOs taken by the Cosmic Origins Spectrograph (COS) on the *Hubble Space Telescope* (*HST*). By correlating the positions of detected absorption lines with the galaxies in their proximity, I aim to gain insights to the complex relationship between intergalactic gas and the galaxies which feed on it. This introduction provides some historical and astrophysical perspective, as well as an overview of the methods and history of QSO absorption line spectroscopy.

Keywords: galaxies:intergalactic medium, galaxies:evolution, galaxies:halos, quasars: absorption lines

1. OVERVIEW OF THE CIRCUMGALACTIC MEDIUM

The majority ($\sim 90\%$) of the baryons in the universe are found *outside* of galaxies, in the diffuse intergalactic and circumgalactic media (IGM/CGM). We define the CGM as the gas for which a nearby galaxy's gravity dominates, which generally extends to approximately a virial radius (see, e.g., Shull (2014)). The IGM, CGM and the galaxies that reside in it are tightly linked by processes such as feedback and accretion. In order to sustain the level of star formation observed, galaxies must accrete gas throughout their lifetimes (e.g., Oort 1966, 1969, 1970; Wakker et al. 1999; Tinsley & Larson 1978; Tinsley & Danly 1980; Erb 2008; Prochaska & Wolfe 2009). At the same time, ongoing star formation and active galactic nuclei (AGN) activity produce feedback that drives gas back into the IGM. This life cycle of gas is complex, and difficult to constrain observationally. Understanding the properties of the IGM, such as its densities, temperatures, motions, and its relationship to the galaxies embedded within it is *essential* for explaining the evolution of galaxies and the star formation history of the Universe.

The properties of the vast reservoir of material in the IGM can be understood by analyzing lines of sight toward background quasi-stellar objects (QSOs). Individual concentrations of gas along a given sightline imprint a ‘forest’ of absorption lines on the spectrum in the direction of the QSO target. The metal lines trace the star formation history within the intervening gas, and neutral hydrogen lines (Ly α is ~~one of~~ the most commonly and easily observed) indicate both the location and velocities of outflowing gas, as well as the presence of fuel

for future star formation. Figure 1 shows an artists impression of a galaxy complete with a CGM halo with arrows indicating outflowing and recycling material, and an illustration of a HST sightline detecting halo gas absorption. The relationship between the galaxies and the IGM is usually studied by looking for galaxies that lie at similar redshifts as detected absorption lines. This approach has value but is incomplete; it does not allow for an unbiased understanding of the distribution of the gas around galaxies, which requires looking for both detections and non-detections of gas, both near as well as far away from galaxies.

The current standard model of structure formation is given by Lambda Cold-Dark-Matter (Λ CDM) cosmology, which predicts the hierarchical growth of large scale structures seeded by initial fluctuations in the dark matter background. In this picture, both galaxies and the IGM should follow the same underlying density profile (e.g., Fukugita & Peebles 2006; Frieman et al. 2008 and references therein). Some observational evidence of this large-scale relationship has appeared recently, such as Wakker et al. (2015), who showed that Ly α absorption strength (equivalent width; EW) traces the overall distribution of galaxies in a Cosmic Web filament. Figure 2 shows their plot of the EW of Ly α absorbers as a function of distance to the center of a galaxy filament, with the enhanced absorption strength evident close to the filament center. In addition, numerous studies have shown that Ly α absorbers also trace individual galaxy halos (e.g., Lanzetta et al. 1995; Chen et al. 1998, 2001; Tripp et al. 1998; Bowen et al. 2002; Côté et al. 2005; Wakker & Savage 2009; Steidel et al. 2010; Prochaska et al. 2011; Thom et al. 2012; Stocke et al. 2013; Tum-

linson et al. 2013; Liang & Chen 2014; Danforth et al. 2016).

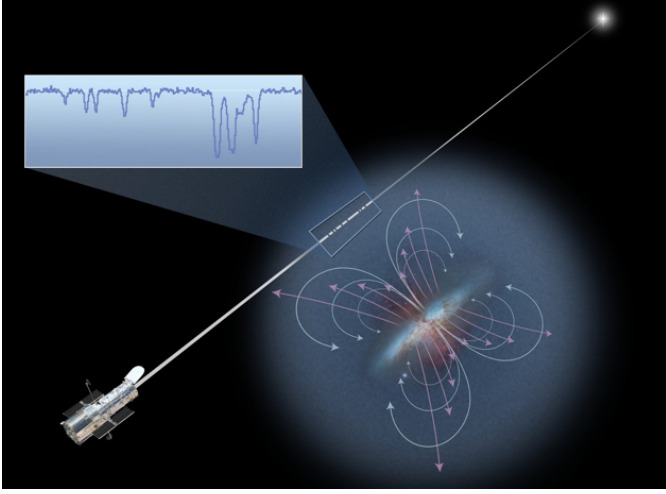


Figure 1. An artist's impression of the CGM of a galaxy. Image credit: NASA/STScI/Ann Feild

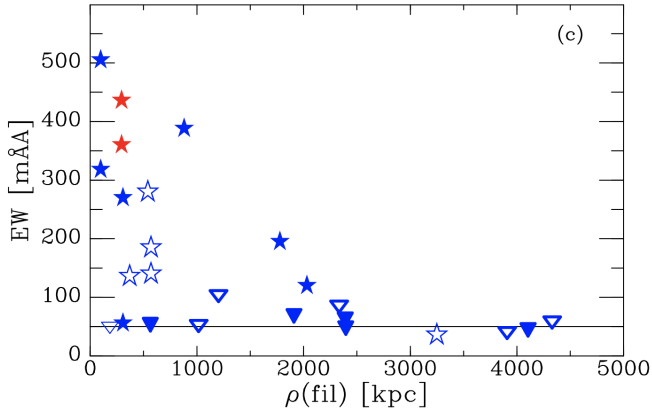


Figure 2. The EW of Ly α absorbers as a function of distance to the center of a galaxy filament. Blue downward tracing triangles indicate upper limits for non-detection, and all stars indicate detections. See Wakker et al. (2015). (the two red stars are for detections within a galaxy's virial radius, the blue stars all occur far from any known galaxy).

The majority of these studies have reported tentative evidence for enhanced Ly α absorption strength with increasing galaxy proximity. The studies of Côté et al. (2005) and Prochaska et al. (2006) found that about half of Ly α absorbers lie within galaxy halos, at impact parameters $\rho \lesssim 350$ kpc. In addition, Wakker & Savage (2009) found that for 90% of $L > 0.1L^*$ galaxies an absorber can be found within 400 kpc and 400 km s $^{-1}$, and all galaxies have a Ly α absorber within 1.5 Mpc. Higher redshift studies, such as Rudie et al. (2012) at $2 < z < 3$, find evidence for an elevated density of absorbers up to 2 Mpc from galaxies. Wakker & Savage (2009) also con-

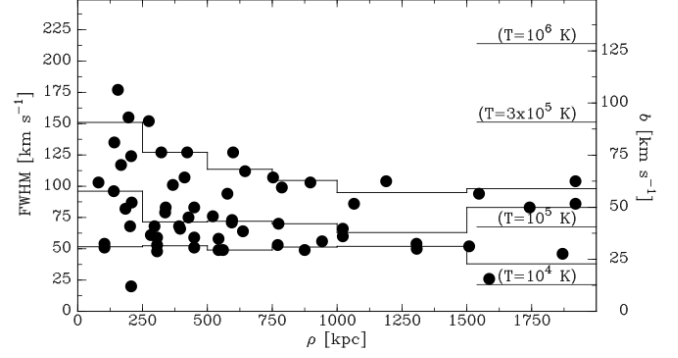


Figure 3. The linewidth (or Doppler b -parameter) and FWHM of Ly α absorbers as a function of physical impact parameter to the nearest galaxy. Histograms show the 10th, 50th, and 90th percentiles of the distribution. See Wakker & Savage (2009).

firmed a previously suggested correlation between Ly α absorption linewidth (also called Doppler b -parameter) and impact parameter (ρ), observing that the broadest lines (FWHM > 150 km s $^{-1}$) are only seen within 350 kpc of a galaxy, while only narrower lines (FWHM < 75 km s $^{-1}$) are found at $\rho > 1$ Mpc (see Figure 3). The more recent COS-Halos survey (Tumlinson et al. 2013 and references therein) studied both the H I and low-to-medium ionization state metals CGM around $\sim L^*$ galaxies, and found that H I is detected nearly ubiquitously within ~ 150 kpc of both star-forming and passive galaxies, with metal absorption lines also detected in the majority of cases but with a stronger dependence on galaxy type (e.g., Tumlinson et al. 2011; Werk et al. 2013).

In addition, studying the enrichment of galaxy halos is necessary for constraining outflow models and informing stellar feedback prescriptions. Directly measuring the velocity field and column densities of absorbers as a function of impact parameter and orientation around galaxies would provide the clearest evidence of inflow or outflow activity, but results are few and uncertain. Kacprzak et al. (2011b) claim to find that Mg II equivalent widths correlate with galaxy inclination but Mathes et al. (2014) find no such correlation for Ly α and O VI absorbers. Furthermore, we should expect outflowing gas to be more highly enriched and trace the metallicity of the associated galaxy, with inflowing gas instead appearing only in H I. Both Stocke et al. (2013) and Liang & Chen (2014) find an “edge” to heavy ion absorption at $\sim 0.5R_{\text{vir}}$, but with Ly α covering fractions of $\sim 0.75 - 1$ continuing out to R_{vir} . However, Mathes et al. (2014) measure O VI absorption out to $\sim 3R_{\text{vir}}$.

All these previous studies have suffered from small sample sizes (most with fewer than 50 systems), and incompleteness due to their higher mean redshifts where it is increasingly difficult to detect faint galaxies sur-

rounding absorption systems. This thesis aims to address some of these issues by compiling the largest-yet survey of Ly α absorbers spread across a range of environments, and both near and far from galaxies. The installation of the Cosmic Origins Spectrograph (COS; Green et al. 2012) on the *Hubble Space Telescope* (HST) has opened up a new era for studying intergalactic gas via UV QSO-absorption lines, as COS is able to observe fainter targets than ever before with high signal to noise and velocity resolution. Over 700 QSOs have now been observed with COS, most of which have good quality data covering the Ly α transition in the nearby Universe ($cz \leq 10,000 \text{ km s}^{-1}$), where the existing galaxy data is also good and relatively complete to $\sim 0.1L^*$ on average.

This project aims to take advantage of all this existing data to study the Ly α traced CGM in the local universe with a survey of unprecedented size. In Chapter 1 I describe the compilation of a new nearby galaxy catalog to take advantage of the existing galaxy data. This catalog is then correlated with the over 700 archival QSO targets observed by the Cosmic Origins Spectrograph (COS) on the Hubble Space Telescope (HST) to produce a sample of galaxy-absorption line systems including a range of galaxy environments, orientations, and types. In the following section I summarize the 3 major questions this thesis aims to tackle with the resulting dataset.

2. SCIENCE GOALS

While there are numerous open questions concerning the interactions of galaxies and the gas that surrounds them, this thesis focuses on the following:

2.1. Galaxy Proximity

How strongly is intergalactic gas concentrated near galaxies, and does the presence of galaxies affect the physical properties of absorbers? Recent studies find that half of all Ly α absorbers lie within galaxy halos, at impact parameters of $\rho < 350 \text{ kpc}$ and within 400 km s^{-1} of a galaxy (e.g., Côté et al. 2005; Prochaska et al. 2006; Wakker & Savage 2009). Furthermore, Sorini et al. (2018) find that the “sphere of influence” of galaxy can extend all the way to $\sim 2 \text{ Mpc}$, far more distant than the $\sim 150 \text{ kpc}$ or $\sim 1R_{\text{vir}}$ often used as the search radius for CGM studies. However, this may just be an effect of galaxies being embedded in filaments, and not necessarily evidence of the influence of the galaxies themselves. The properties of the lines also appear to change with impact parameter. For example, it has been known for some time that higher column density absorption is found closer to galaxies, and there is good evidence that the same is true for absorption linewidth ~~and FWHM~~ (e.g., Wakker & Savage 2009; Prochaska et al. 2011). However, it remains unclear which physical process is responsible; increased turbulence, temperature, ionizing radiation field, or an effect of velocity gradients or blending of multiple cloudlets along the line of sight. Studying this phenomenon as a function of environment with

good understanding of the properties, morphologies, and group memberships of the nearby galaxies is the clearest way forward on this question, and a larger statistical sample than previously employed will be required.

2.2. Galaxy Orientation

Do the physical properties of absorbers depend on their orientation with respect to nearby galaxies? Recent results suggest that absorbing systems have a preferred orientation with respect to the major and minor axes of the galaxies they are near to (e.g., Kacprzak et al. 2011b, 2012). This could be evidence of inflows and outflows, an effect of the global structure of galaxy halos, or a signature of a preferred orientation of galaxy halos within Cosmic Web filaments. Unfortunately the statistics are not yet good enough to distinguish between these possible scenarios. Additionally, very few authors (see, e.g., Mathes et al. (2014); Bordoloi et al. (2014)) have investigated the dependence of absorber properties on nearby galaxy *inclination*, the results of which could have important implications for galaxy halo shape and the spatial dependence of halo gas covering fractions. A large-scale study into the inclination and azimuth angle dependence of absorption is the clearest path forward here.

2.3. Galaxy Rotation

Does intergalactic gas “know” about the rotation of the galaxies embedded within it? In particular we would like to know how far out (or to what impact parameter) the rotational curves of galaxies extend, or in other words what angular momentum information is retained by galaxy halos. Galaxy disks are built via accretion of material from the IGM, which carries with it angular momentum. This angular momentum must eventually contribute to the disk rotation, so it is reasonable to expect the overall rotation signature of halo gas to trace that of the more readily measured disk gas rotation. Indeed, the simulations of Stewart et al. (2011a,b, 2013) predict that H I gas out to at least $1R_{\text{vir}}$ should co-rotate with galaxies, and furthermore that absorption lines in QSO sightlines should be able to accurately trace this (see Figure 4; a simulated galaxy halo showing coherent halo gas rotation from Stewart et al. 2011a). Previous studies (e.g., Steidel et al. 2002; Côté et al. 2005; Wakker & Savage 2009; Kacprzak et al. 2011a) were unable to find a clear correlation between the rotation of galaxy disks and the kinematics of nearby absorbers. However, none of these previous studies have been able to produce a sample of more than a handful of systems, and have only considered the possibility of an extended, warped stellar disk in their analysis. A targeted observational campaign for easily observed galaxies in the local Universe could make substantial progress here.

3. SUMMARY OF THESIS

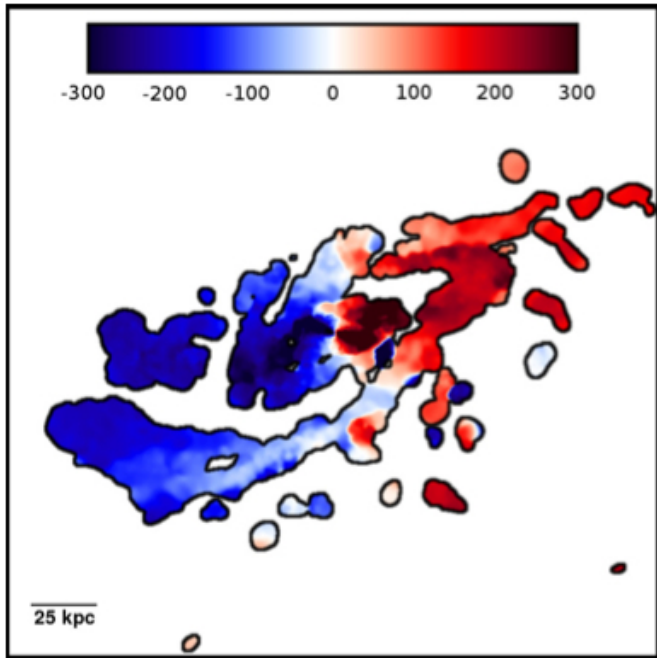


Figure 4. A simulated galaxy halo showing the coherent co-rotation of halo gas out at least 100 kpc from the inner disk (the galaxy disk is seen in dark blue and red at the image center; see [Stewart et al. \(2011a\)](#)).

In the following chapters I describe a program to observationally explore the connections between low column density Ly α absorption and the galaxy environment in the local Universe. This program focuses mainly on archival QSO observations taken by the Cosmic Origins Spectrograph (COS) on the Hubble Space Telescope (HST), and correlations between the locations of detected Ly α absorption and of galaxies larger than $\sim 0.1L^*$. By restricting our study to low-redshifts ($cz \leq 10,000 \text{ km s}^{-1}$) we are able to compile a dataset of unparalleled size while remaining highly complete to galaxies of all types, sizes, and distances from absorption detections. The results of this thesis are presented as follows:

1. In Chapter 1 I present a new nearby galaxy catalog. In order to study the CGM-galaxy connection on a large, all-sky scale, we rely heavily on archival, publicly available data for the positions and properties of the galaxies. We describe the retrieval, handling, homogenization and completeness of these data, as well as detailed descriptions of each included galaxy property (i.e., the catalog columns).

2. In Chapter 2 I present the results of a pilot study with 33 QSO sightlines chosen for their proximity to large galaxies ($D \geq 25 \text{ kpc}$). We introduce a new method for absorber-galaxy matching called the likelihood-method, which will make it possible to algorithmically study our large final data set. Using this likelihood-method we match 48 Ly α absorption lines with nearby large galaxies, and study the absorption strength (EW) as a function of velocity and spatial separation, azimuth angle and inclination. We find that the strongest absorbers are all found within 100 km s^{-1} of a galaxy, and that there exists an overabundance of detections near highly inclined galaxies ($inc \gtrsim 50^\circ$). We attribute this overabundance to the effect of flattened, non-spherical galaxy halos on the detection probability as a function of impact parameter.

3. In Chapter 3 I present the results of a study of the kinematic connection between galaxy disks and Ly α -traced halo gas. We have compiled a sample of 29 galaxies with known rotation curves both from the literature and from new observations with the Southern African Large Telescope (SALT) which also appear within $3R_{\text{vir}}$ of a COS QSO sightline. We compare the galaxy disk kinematics to the velocities of Ly α absorption lines detected in 19 nearby QSO sightlines with the help of custom cylindrical and NFW-based halo rotation models ([Navarro et al. 1996, 1997](#)). We find that the co-rotation fraction of absorbers declines as a function of galaxy luminosity, which we attribute to the effect of cold-mode accretion dominating in lower-mass galaxy halos.

4. In Chapter 4 I present the results of our full CGM survey, which includes 1135 Ly α absorbers detected in the spectra of 264 QSO spectra. We explore the effect of different normalizations to our likelihood-method for absorber-galaxy matching, and use this technique to split absorber-galaxy systems into 5 different bins based on their galaxy environments. We find that both absorption strength (EW) and linewidth (Doppler b -parameter) are enhanced with proximity to a single galaxy, and further enhanced by proximity to multiple galaxies. We also detect a bimodal azimuth distribution, with Ly α absorbers preferentially found slightly offset from both the minor and major galaxy axes. We confirm the inclination results first suggested in Chapter 2.

5. In Chapter 5 I summarize the results of this thesis, and place these results in the broader context of the circumgalactic medium and its implications for global galaxy evolution.

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